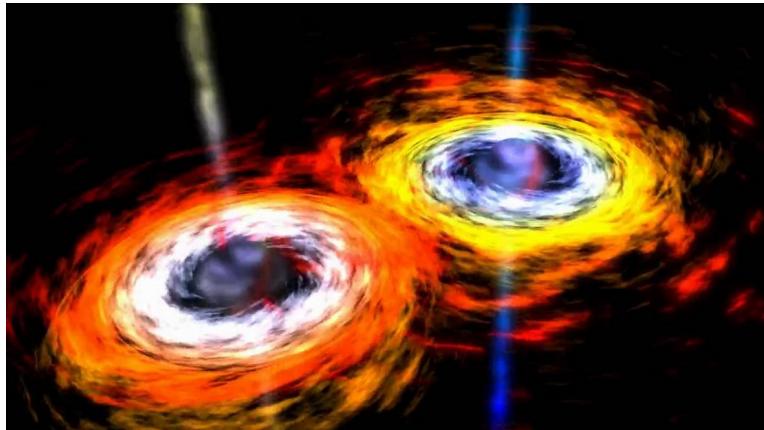
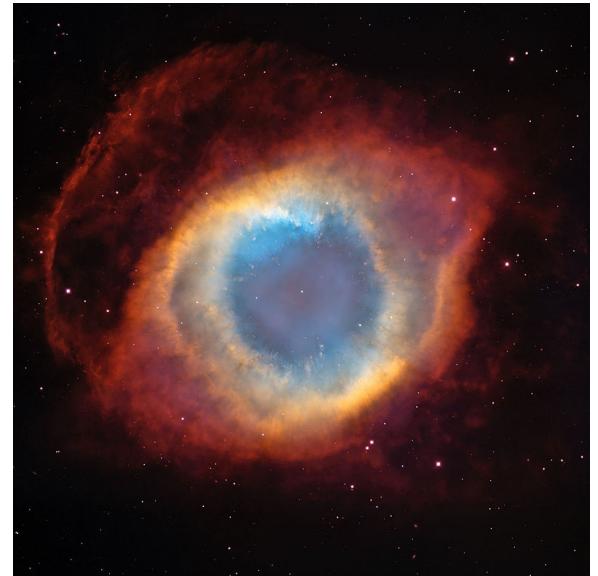


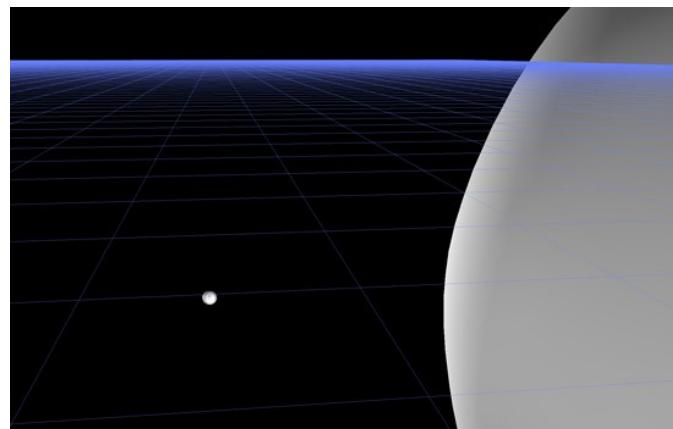
Gravitational Waves – Sources and Detectors



Orbiting Massive Bodies



Supernova, NGC7293-Helix Nebula



Extreme Mass Ratio Blackhole Mergers

Credit: <https://www.rit.edu/news/story.php?id=47936>

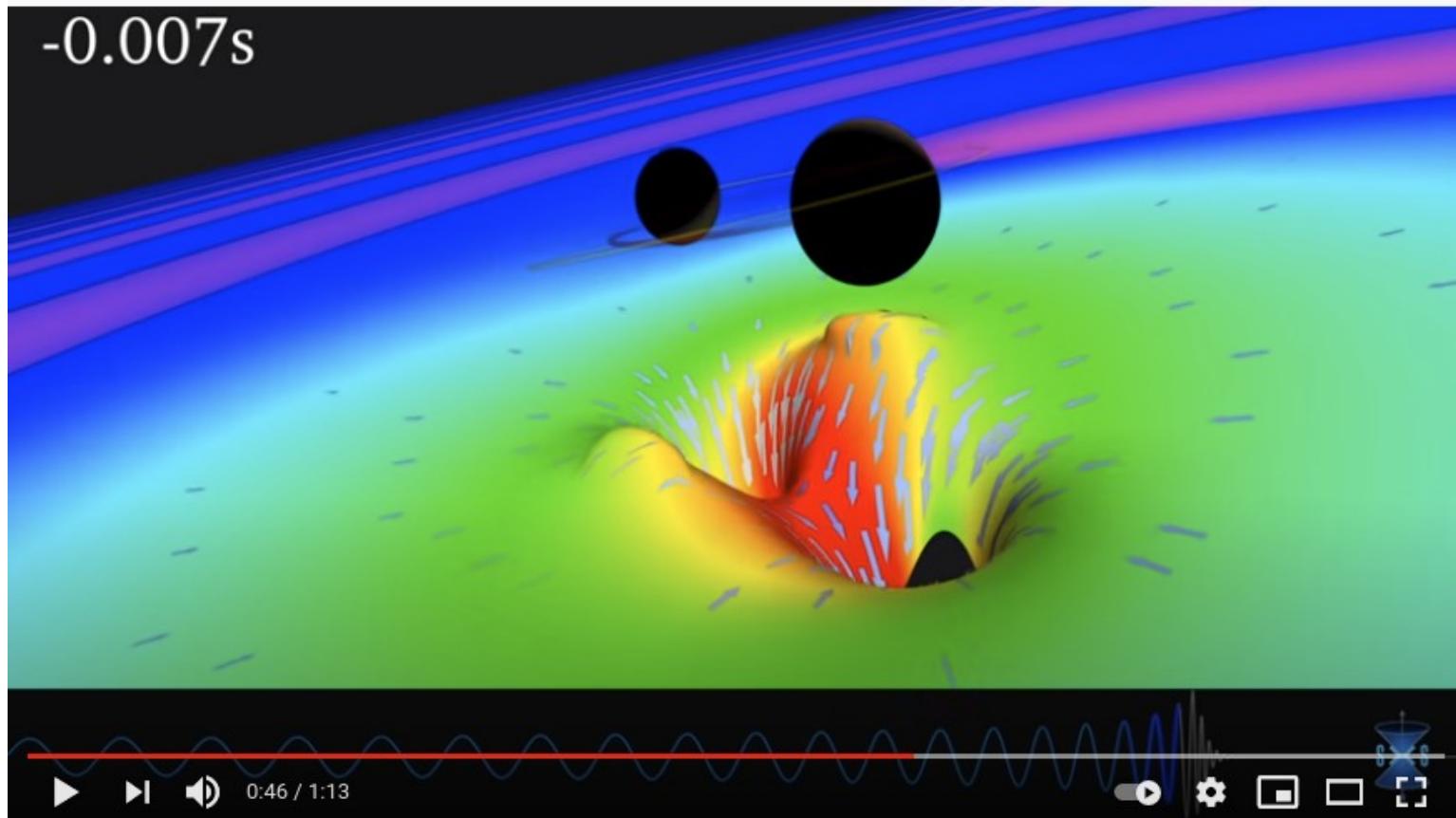


Sources of Gravitational Waves

- **Ripples in space time**...hard to do, spacetime is very “rigid” or “stiff.”
- Need **matter in motion** to create Gravitational Waves, recall charges in motion create electromagnetic waves.
- The motion **cannot be spherically symmetric**. Conservation of mass and conservation of linear momentum (of masses) means the first “moment of the mass distribution” that can cause GWs is “quadrupole.”
- **Large masses and fast changes in time** create stronger GWs.
 - Black holes are small and massive. In binary orbits with another black hole, orbital speeds can reach more than 30% the speed-of-light! Small is good because they can get very close before they “touch” or “merge.”
- **Sooooooo....**

Sources, Binary Black Holes

- **Binary Black Holes** inspiral, merger, and ringdown. First discovered/seen **GW150914**, announced April 2016.



<https://youtu.be/1agm33iEAuo> 1m13s, no matter around, GW150914 parameters
Gabella GravWaves



Sources, Binary Neutron Stars

- **Binary Neutron Stars** also inspiral, merger, and ringdown. First **GW170817**. Seen in all optical and other EM bands.

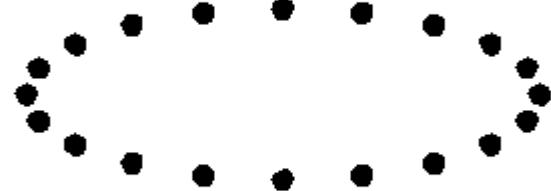


<https://youtu.be/1agm33iEAuo> 48s, matter around, kilonova, heavy elements created.

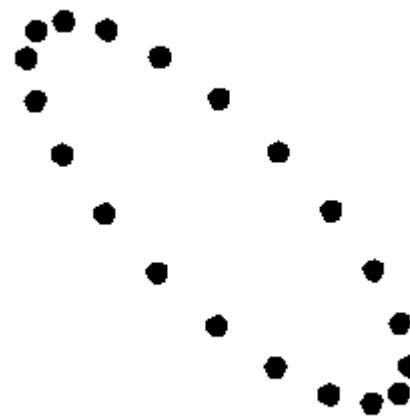
Gabella GravWaves

Nature of Gravitational Waves, and $h_{\mu\nu}$

Of all 10 possible components to h , it simplifies to two polarizations. One with a motion that shrinks (grows) in one direction while growing (shrinking) in the other direction. The other polarization is just rotated by 45 degrees.

$$h_+$$


hplus polarization

$$h_\times$$


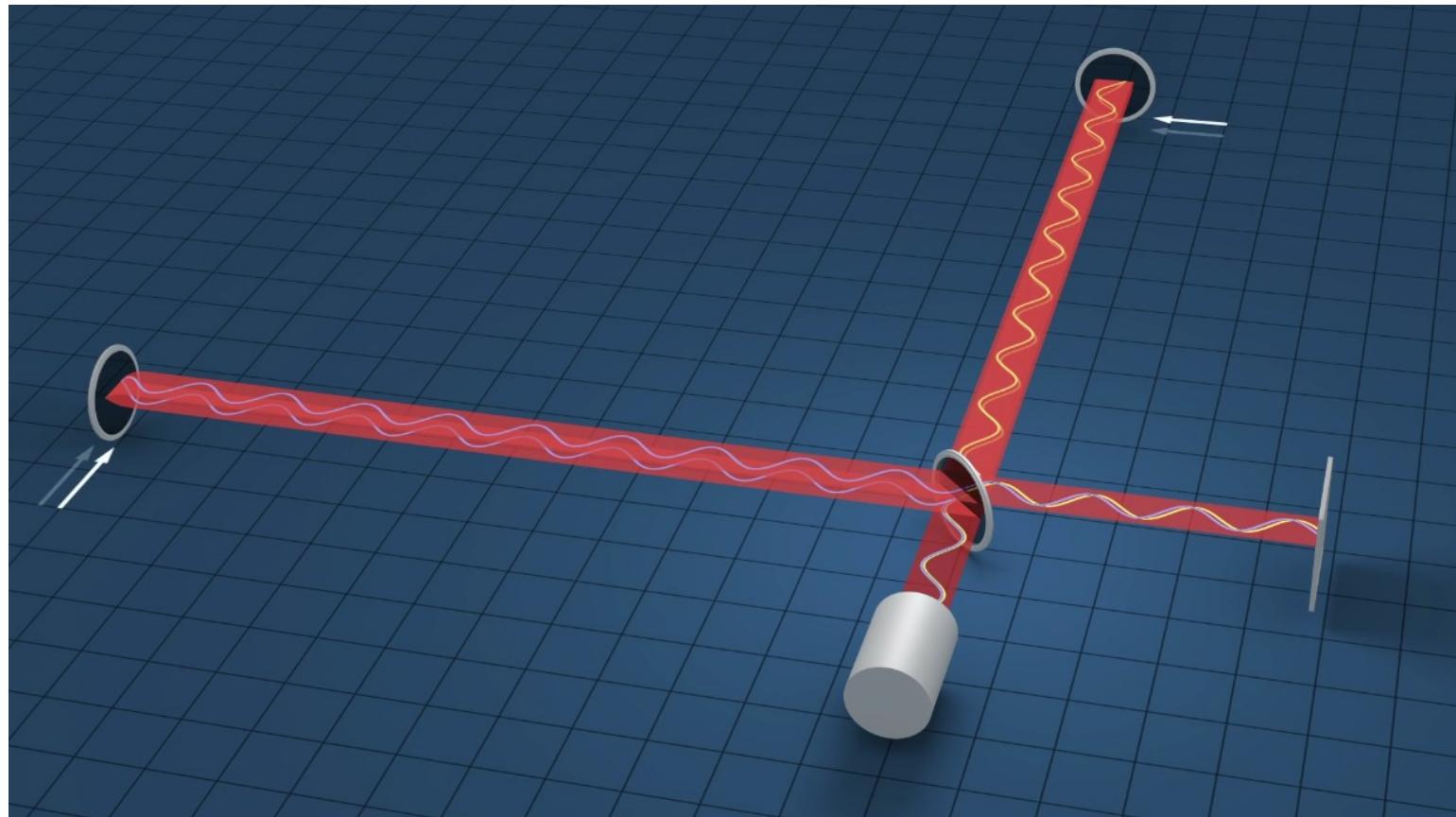
hcross polarization

Wave headed into the page, Y up, X left.

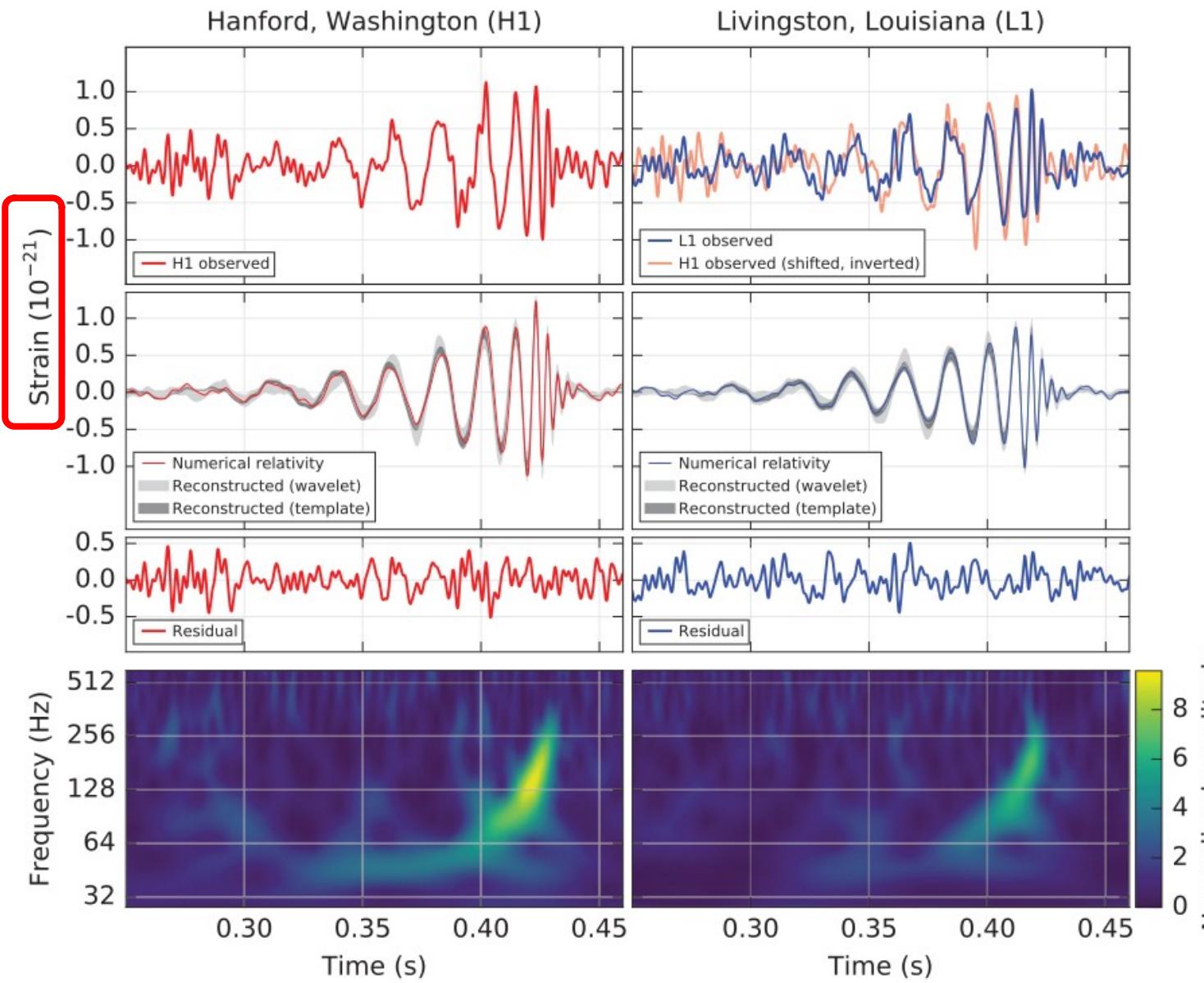
Gabella GravWaves

Big Michelson Interferometer, 4 km arms

- YouTube link https://youtu.be/tQ_teIUb3tE



Gravitational Wave GW150914



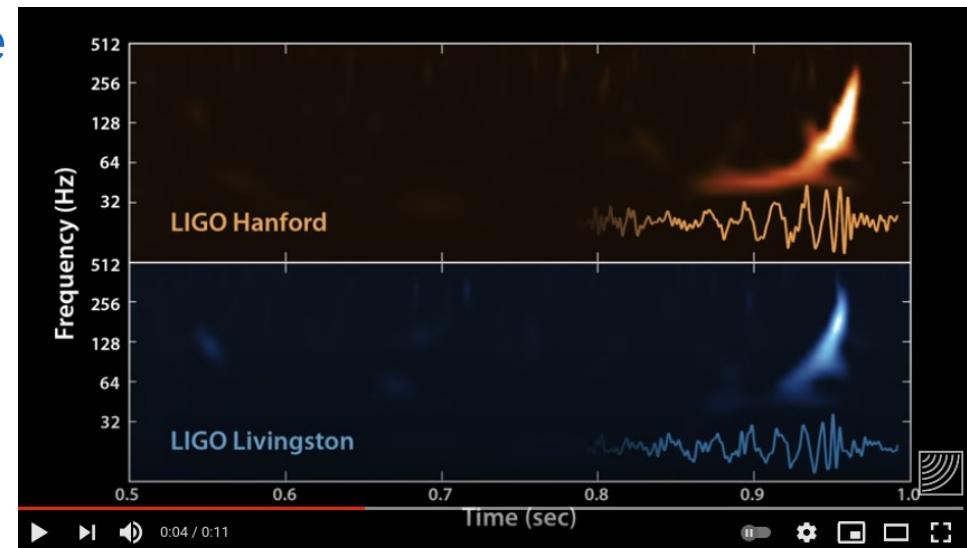
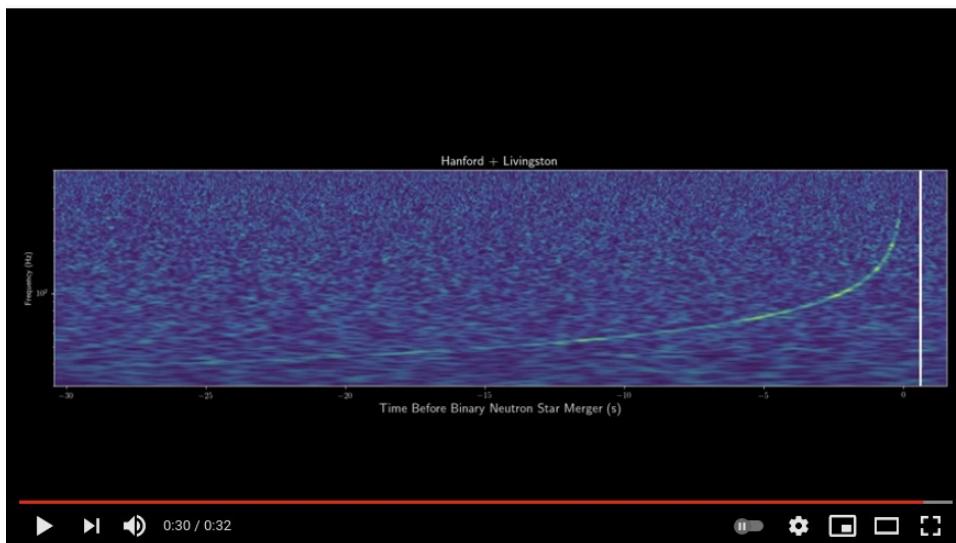
Direct Discovery by LIGO of a gravitational wave announced on 11 February 2016 for a wave on 14 September 2015---event labeled GW150914.

SNR = 24, 5 sigma detection!

Pretty after filtering out the known seismic frequencies.

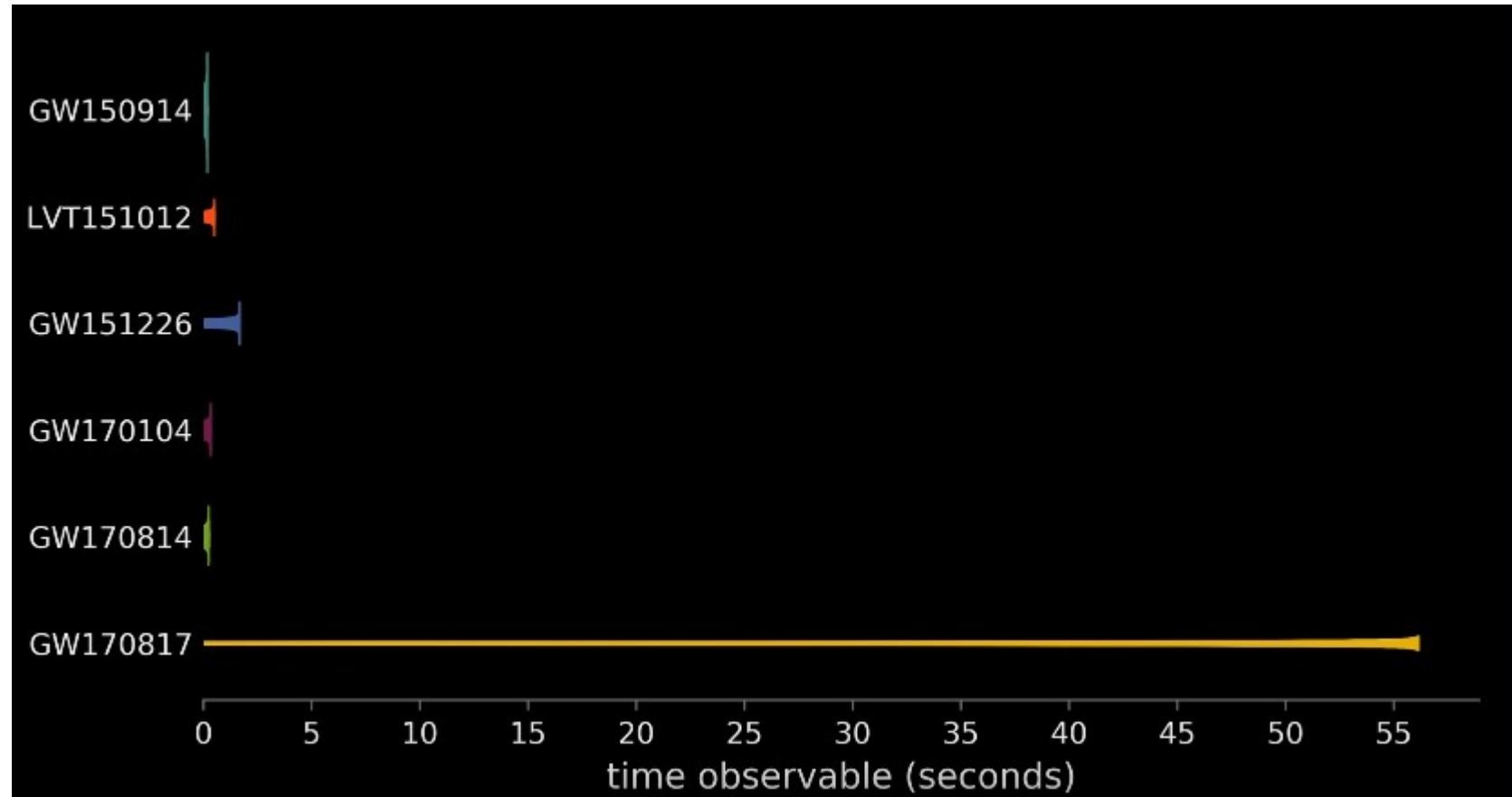
Sources, Waveforms---The Strain

- **Waveforms**, the strain in the interferometer(s)
- Binary Black Holes, GW150914, [here](#)



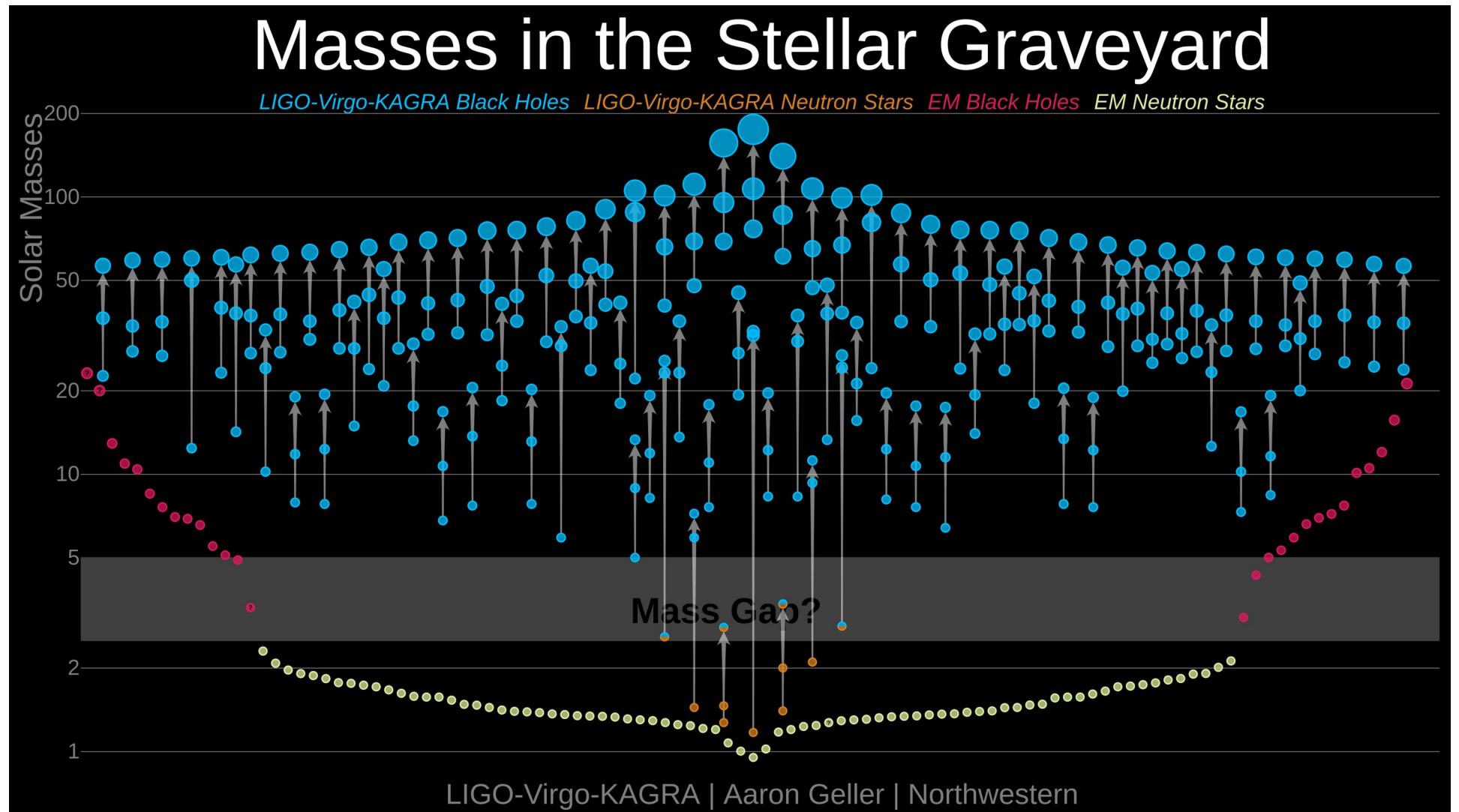
- Binary Neutron Stars, GW170817, [here](#)

GW170817 with other BH GW Sources



YouTube video, animated waveforms, [here](#)

Population of Black Holes and Neutron Stars



<https://media.ligo.northwestern.edu/>

Reference to 2020 Image

Gabella GravWaves

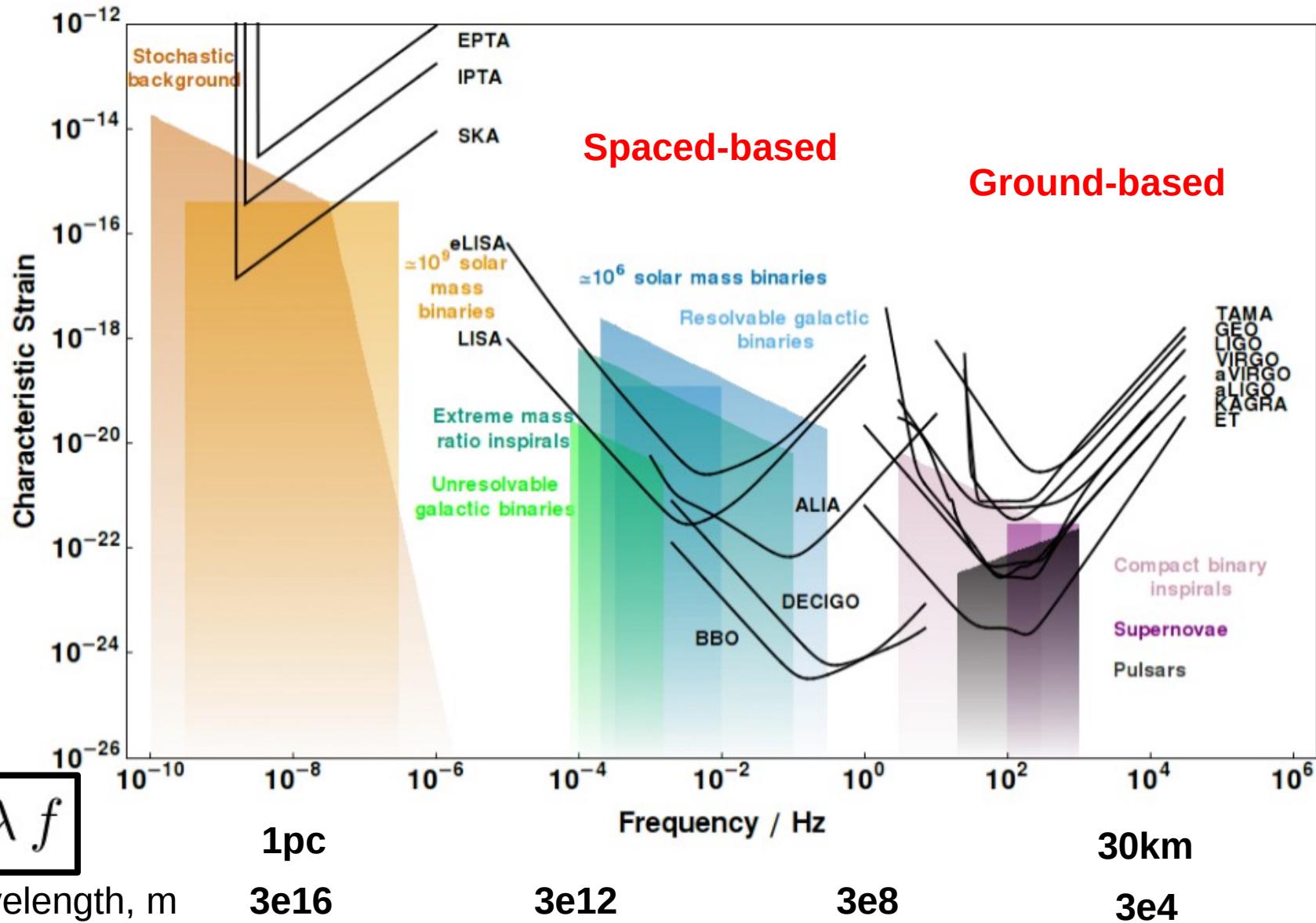
above in 2022

10

Strain Curves from Moore et al.

Pulsar Timing Arrays

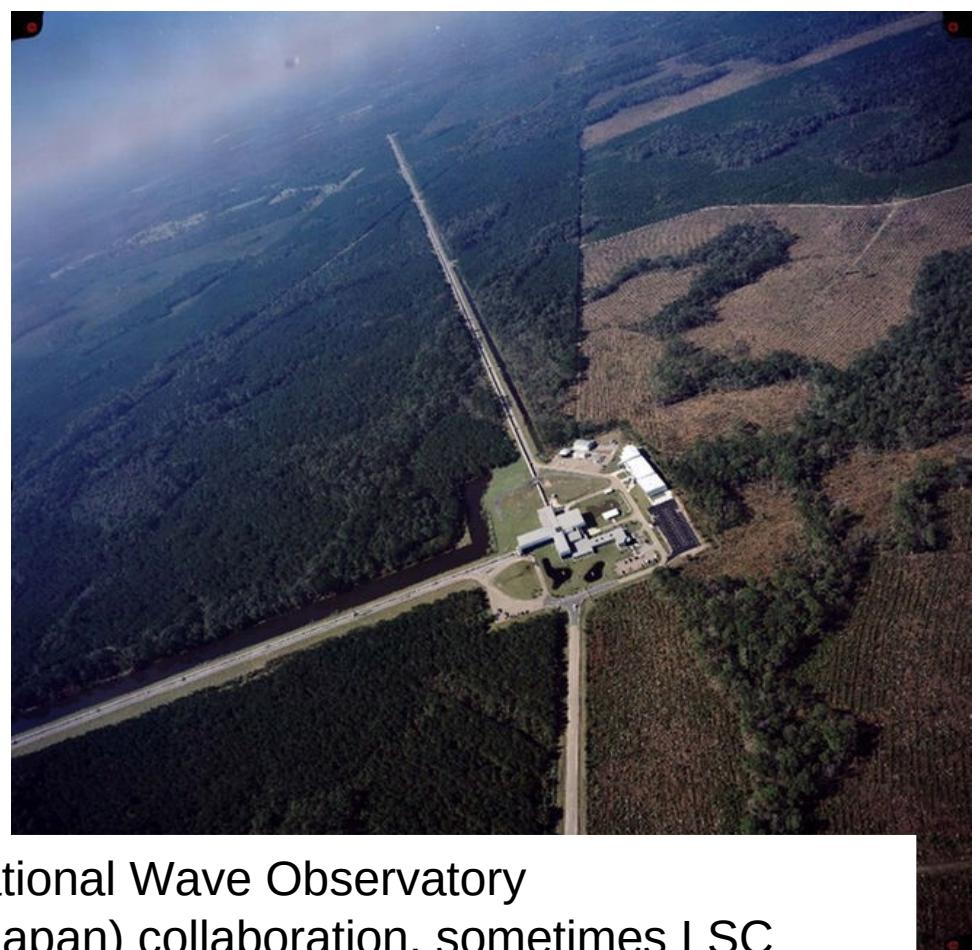
<http://gwplotter.com/>



LIGO, ground-based Gravitational Wave Detectors



LIGO Hanford, WA (LHO)

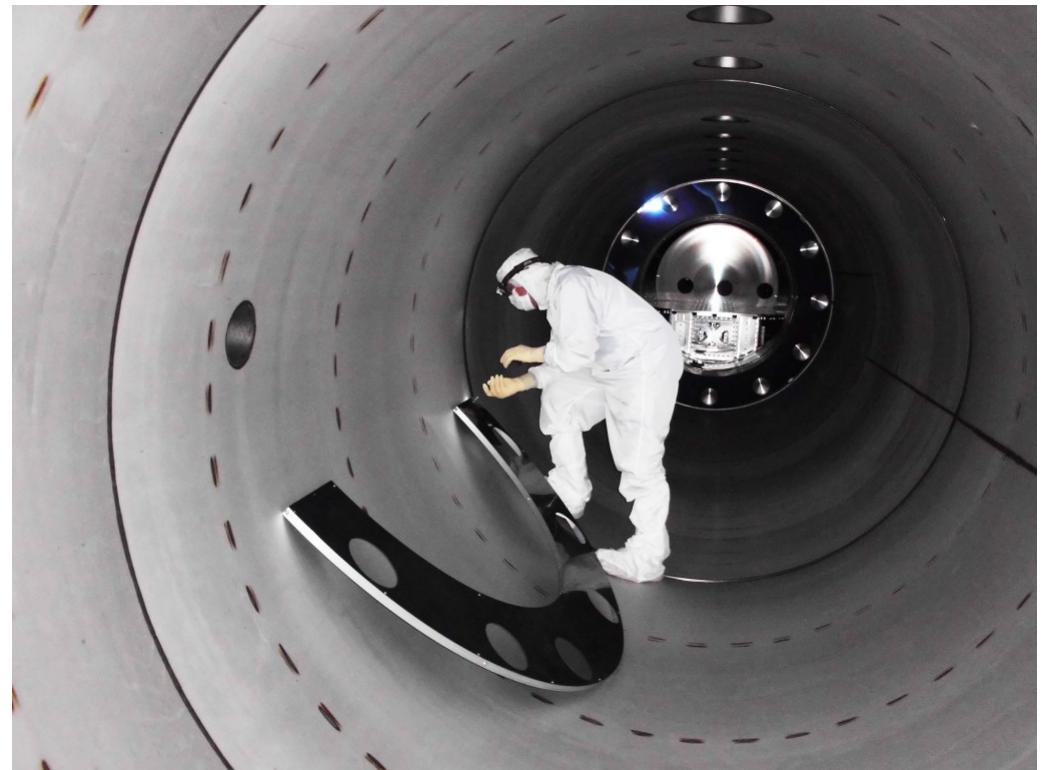


LIGO Livingston, LA (LLO)

LIGO – Laser Interferometric Gravitational Wave Observatory
LVK – LIGO, Virgo (Italy), KAGRA (Japan) collaboration, sometimes LSC
Gabella GravWaves

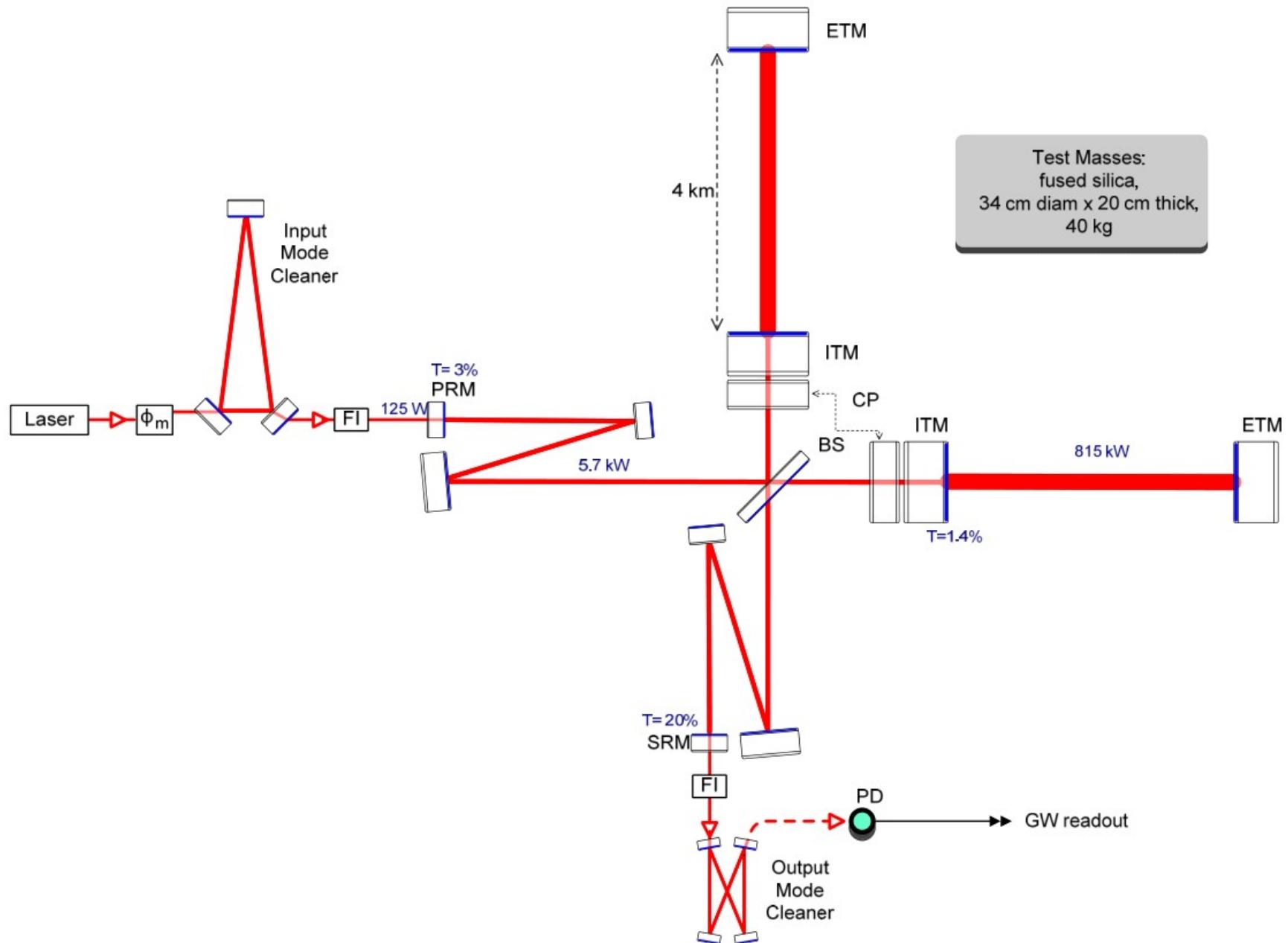
LIGO Hardware

- LIGO “Test Mass” in 4-element suspension



- LIGO Vacuum Tube

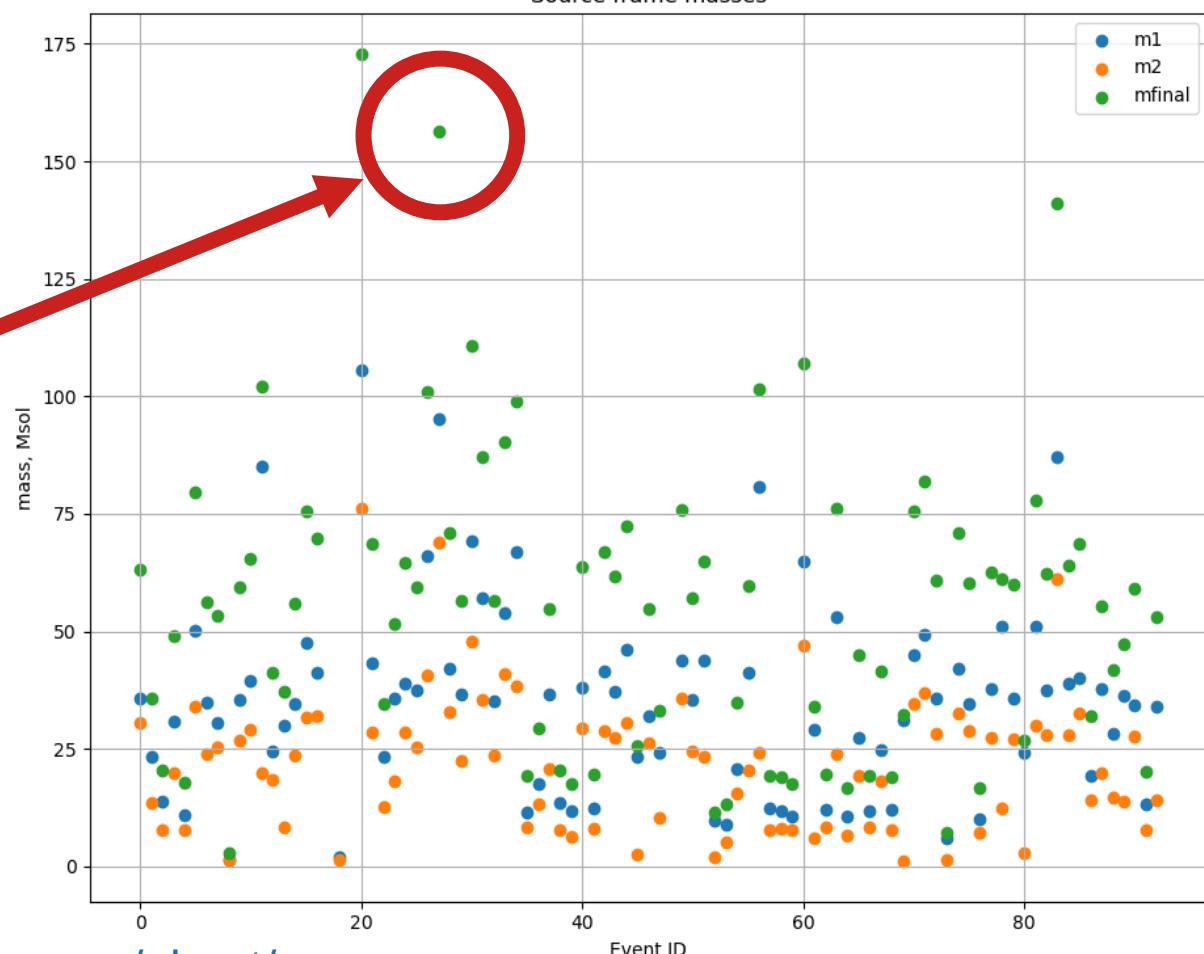
Ref: some hardware [images](#)



Catalog GWTC-3

The publicly released table of gravitational wave events, see the page
<https://www.gw-openscience.org/eventapi/html/GWTC/>

Event **GW190521_v3**, first
Intermediate Mass Black
Hole (IMBH), 156.3 Msol



<https://www.gw-openscience.org/about/>



LIGO Sites

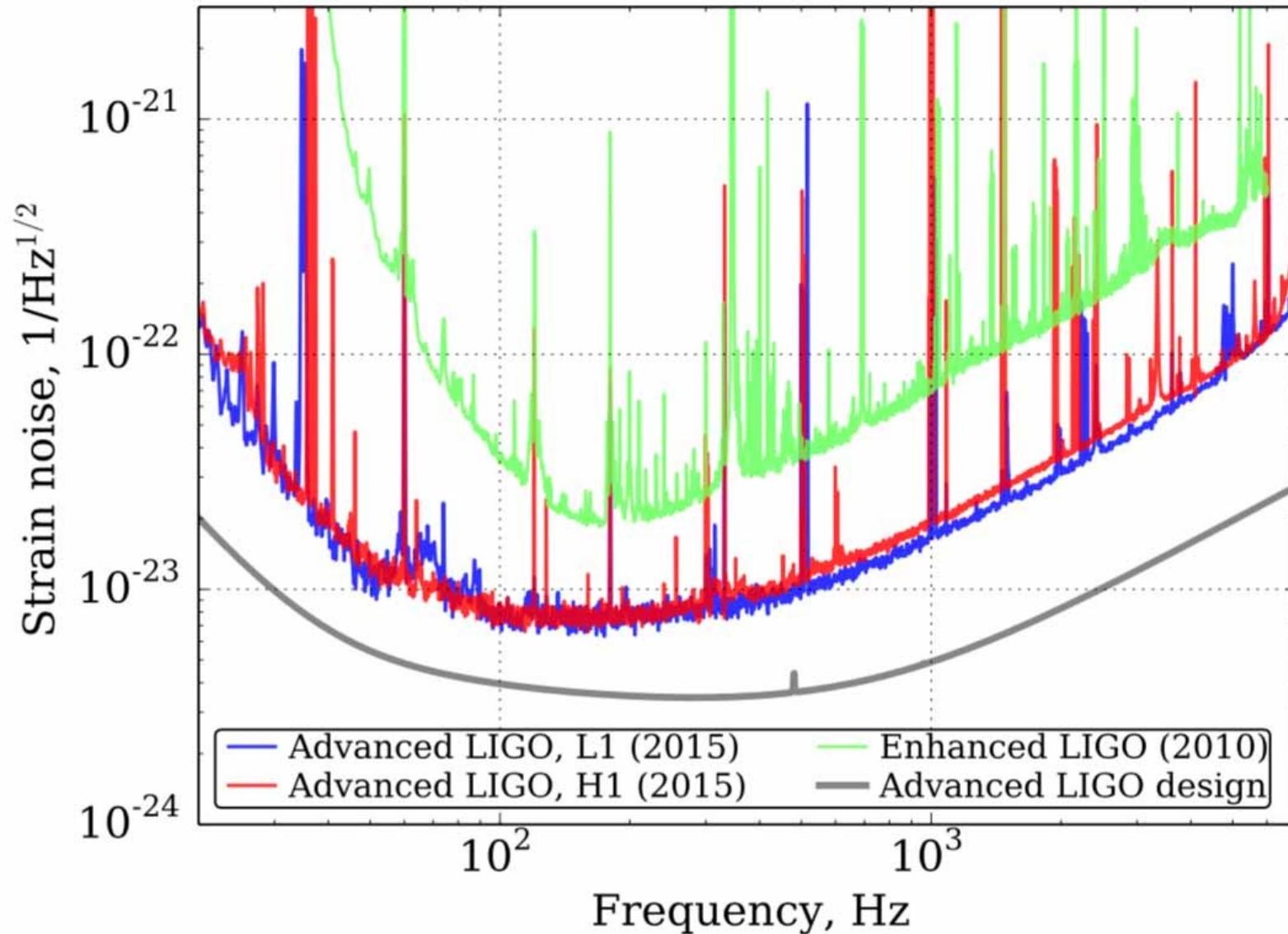
*Hanford
Observatory*



LIGO-G9900XX-00-M

Also [Virgo](#), [Cascina](#), Italy and [KAGRA](#), Japan

Seismic Noise aLIGO and initial



Summary

- Vanderbilt University (Kelly Holley-Bockelmann, Karan Jani, and group) are in LIGO, since September 2020.
 - Led by Karan Jani, newly hired assistant professor at Vanderbilt, researching Intermediate Mass Black Holes and Lunar based GW detectors.
- LIGO should start taking data again March 2023.
 - Shutdown at the beginning of the pandemic, March 2020.



Thank you for your
Attention!

Links

- LIGO Teacher's page [here](#)
- LIGO Student's page [here](#)
- LIGO CalTech Educational Resources [here](#)
- LIGO Scientific Collaboration, [LSC](#)
- LIGO [Document Server](#)
- CalTech GW media assets [page](#).
- Kelly Holley-Bockelman's TEDx Nashville [Talk](#) 2016 about GW150914.
- Pulsar timing array, [Nanograv](#)

Backup

Saulson, If Light Waves...

- Argues that in the expanding Universe, light continuously experiences the increase in scale by expanding continuously itself.

$$\frac{\lambda_1}{\lambda_0} = \frac{R(t_1)}{R(t_0)}$$

- **NOT** the case for a light wave in LIGO...

- The wave filling the cavity **IS** stretched by the sudden onset of the gravitational wave.
- Light travel time, $2 \times 4\text{ km}$ is $27\mu\text{s}$, but GW oscillates at 20 Hz , 50 ms .

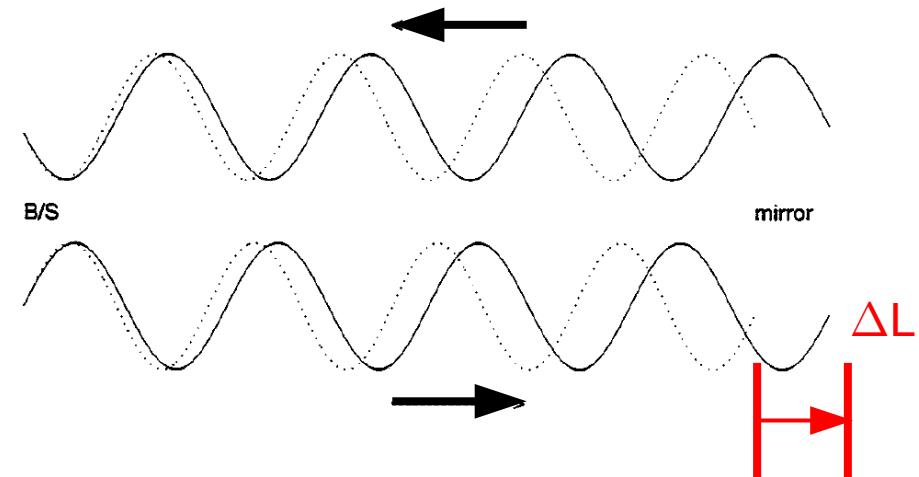
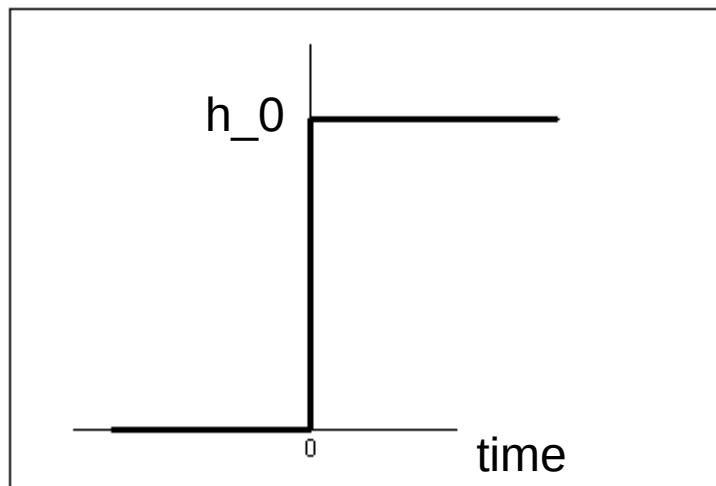


Fig. 2. Light before (dotted) and after (solid) the arrival of a gravitational wave. The beamsplitter is at left, end mirror at right. Outbound light is shown at the bottom, returning light at the top.

Ref: Saulson, *If Light Waves...*, AmJPhys **65**, 501 (1997)
Gabella GravWaves

As the Wave runs through...

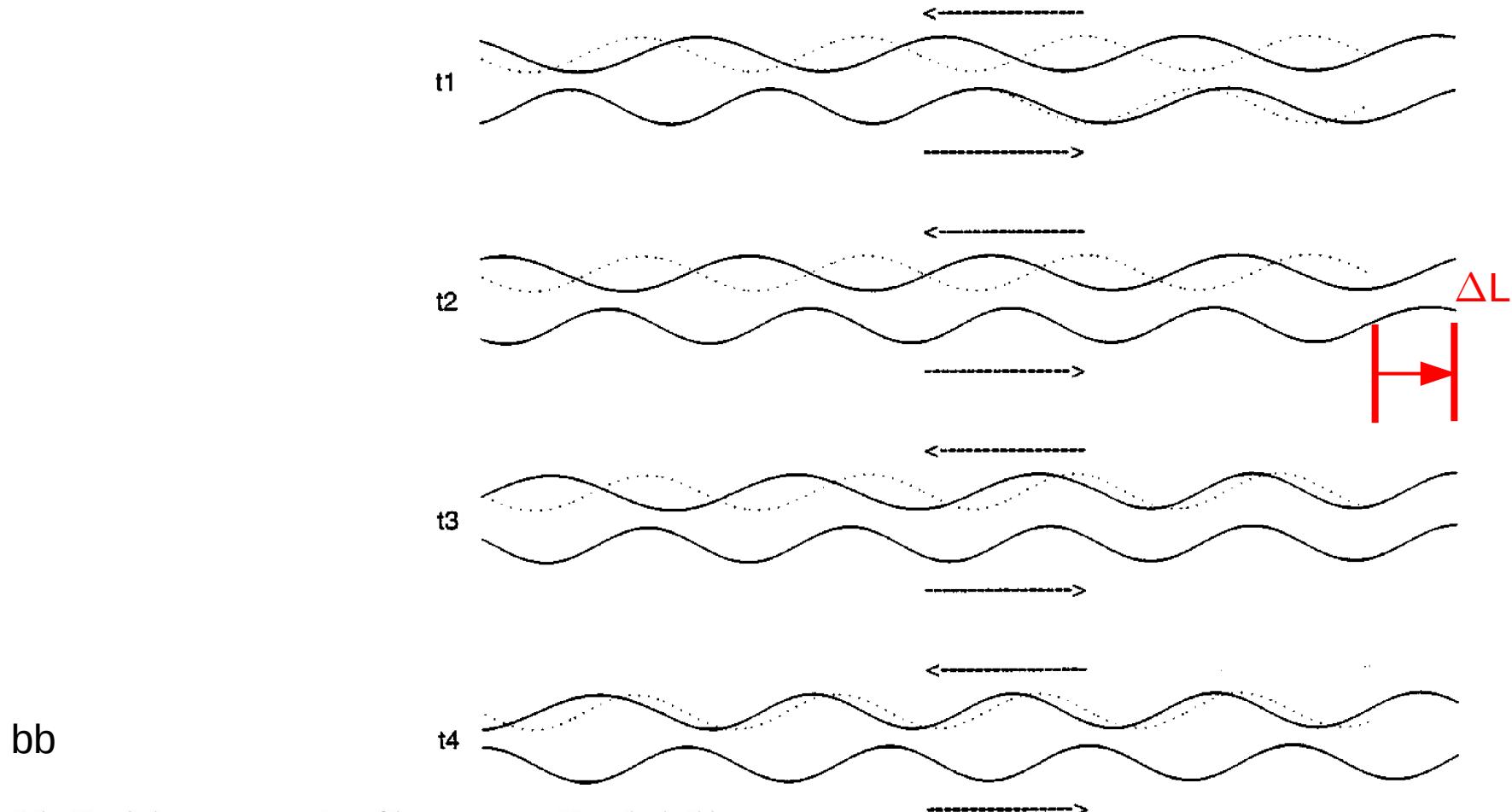


Fig. 3. Like Fig. 2, but at a succession of later moments. Note the buildup of phase shift between the light in the stretched arm (solid) compared to how it would have traveled through an unstretched arm.



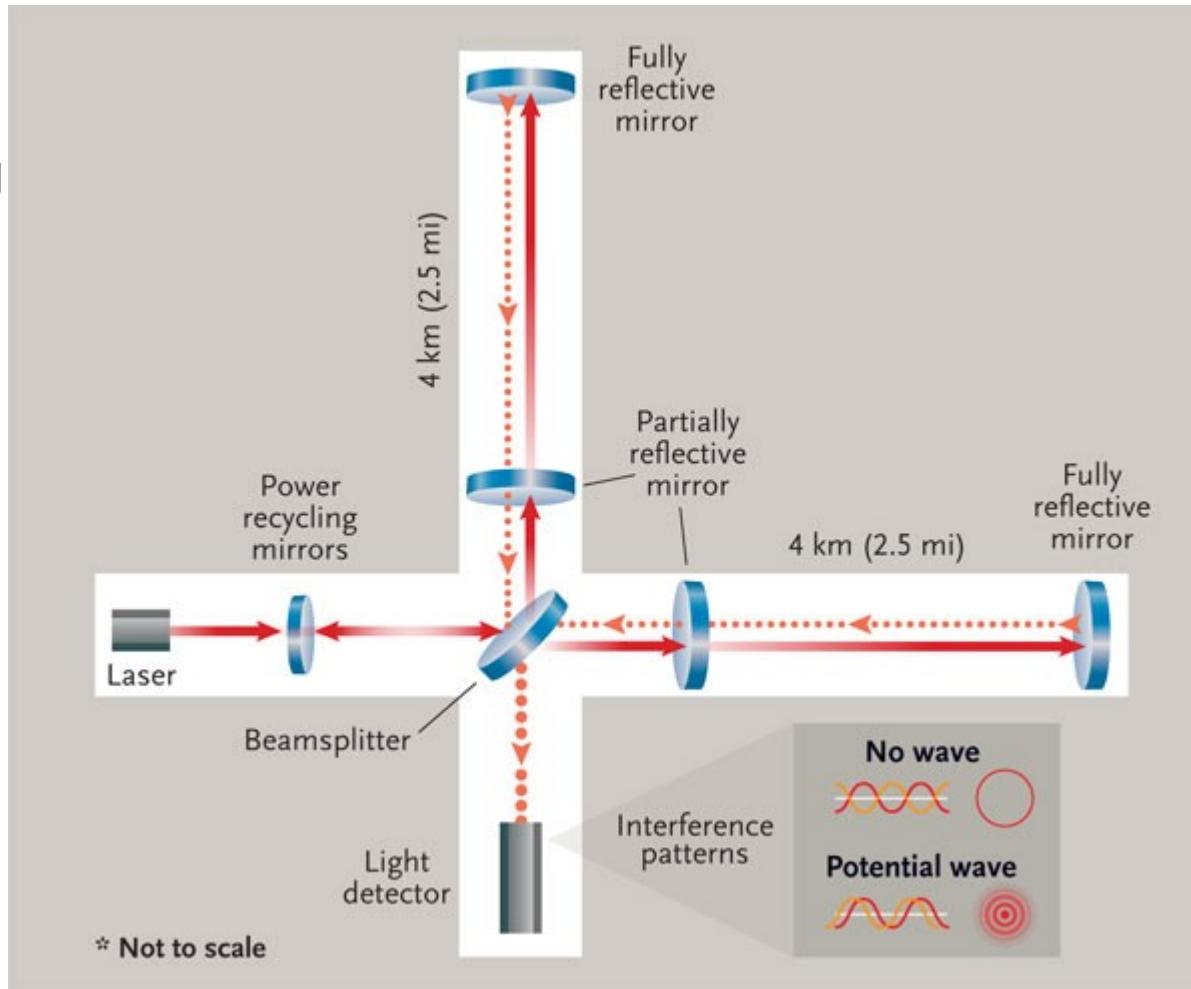
Saulson, If Light Waves..., says rods are still rigid!

V. LENGTHS IN COSMOLOGY AND IN LABORATORY PHYSICS

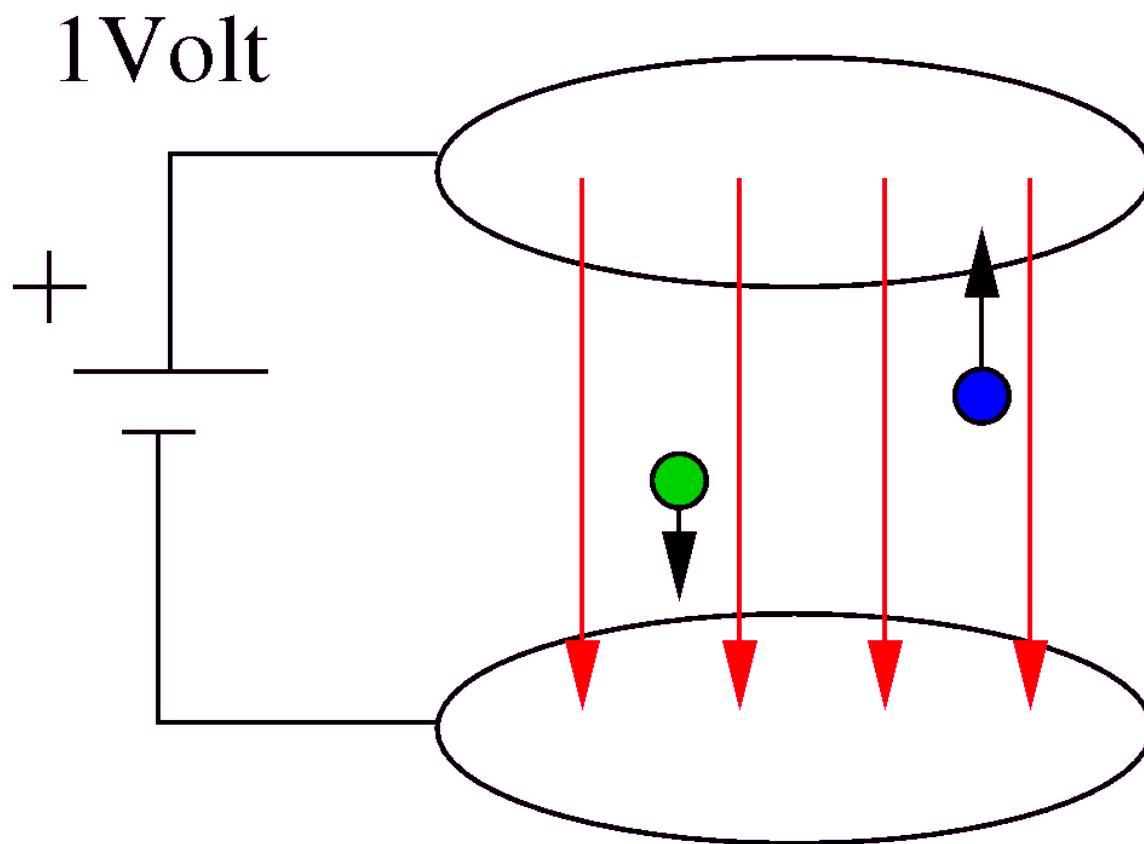
Note that the language we have been using in this paper only makes sense if we imagine that we have standards of length other than either the separations of freely falling test masses or the wavelengths of light waves. We do. A good paradigm of a length standard is a perfectly rigid rod. Such a rod does not change its length in the presence of a gravitational wave, because the arbitrarily strong elastic forces between its parts resist the gravitational force carried by the gravitational wave. As we will see below, we can also use the travel time of light as a reliable ruler under most conditions, in spite of the stretching of light waves that goes on when space expands.

LIGO Interferometer

- 4km long



Units?



- Proton,
heavy, +e
- Electron,
light, -e



SI Prefixes

Table 5. SI prefixes

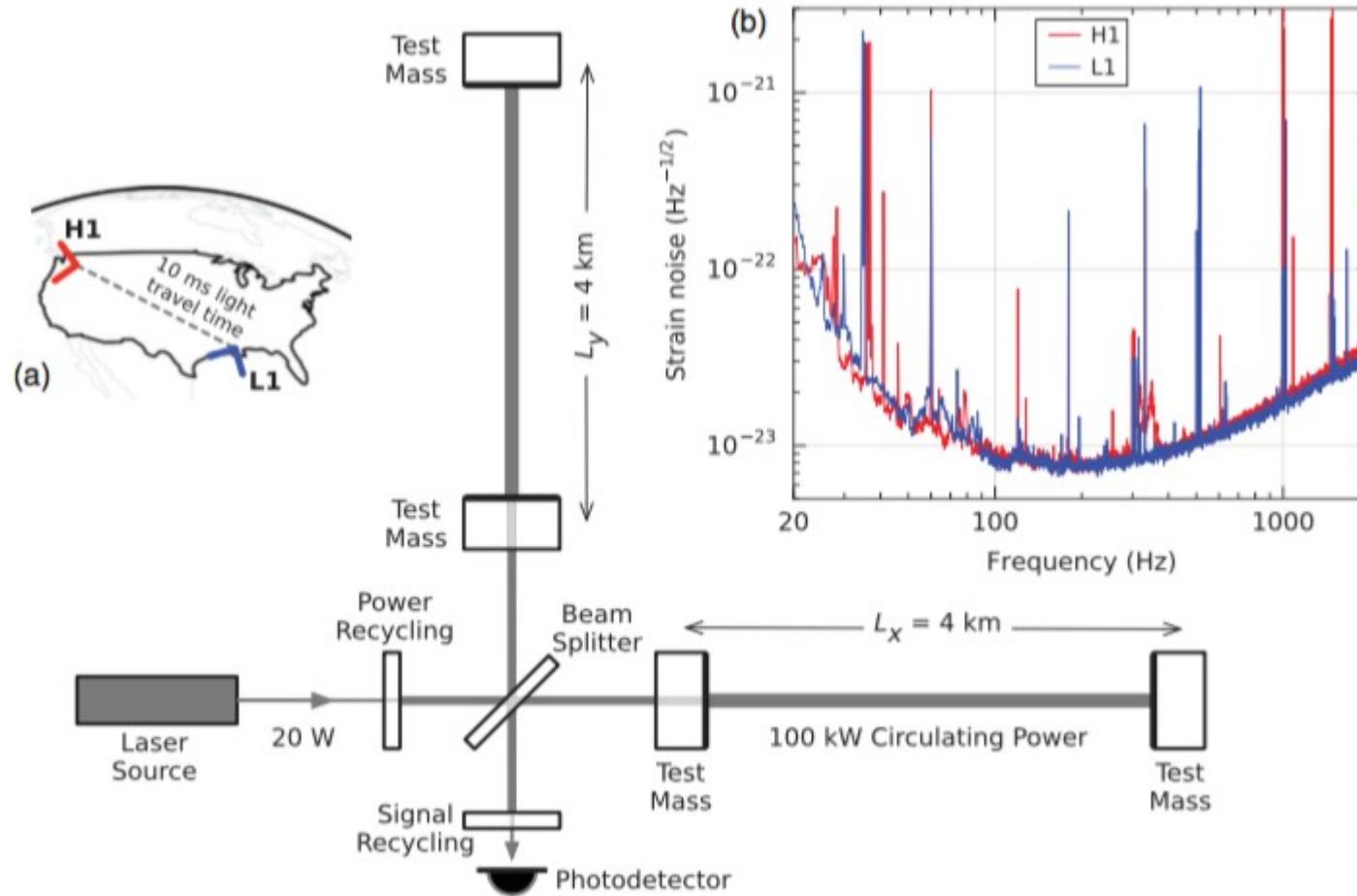
Factor	Name	Symbol	Factor	Name	Symbol
10^{24}	yotta	Y	10^{-1}	deci	d
10^{21}	zetta	Z	10^{-2}	centi	c
10^{18}	exa	E	10^{-3}	milli	m
10^{15}	peta	P	10^{-6}	micro	μ
10^{12}	tera	T	10^{-9}	nano	n
10^9	giga	G	10^{-12}	pico	p
10^6	mega	M	10^{-15}	femto	f
10^3	kilo	k	10^{-18}	atto	a
10^2	hecto	h	10^{-21}	zepto	z
10^1	deka	da	10^{-24}	yocto	y

Merging Black Holes



ref: <http://www.techinsider.io/binary-black-holes-confirmed-gravitational-waves-2016-2>
Gabella GravWaves

Schematic, Map, and Noise



GW strain for Circular Orbit

$$h_0 = \frac{r_{s1} \cdot r_{s2}}{r \cdot R}$$

$$\frac{\omega_s^2}{c^2} = \frac{(r_{s1} + r_{s2})}{2 R^3}$$

$$h_+(t) = h_o \left(\frac{1 + \cos^2 i}{2} \right) \cos 2\omega_s t$$

$$h_\times(t) = h_o \cos i \sin 2\omega_s t$$

- rs1 is $2GM_1/c^2$, Schwarzschild radius for mass M_1 , etc.
- r is distance from Earth to system.
- R is the separation of the two bodies.

Gabella GravWaves

ref: Maggiore around Eqn. 3.332.

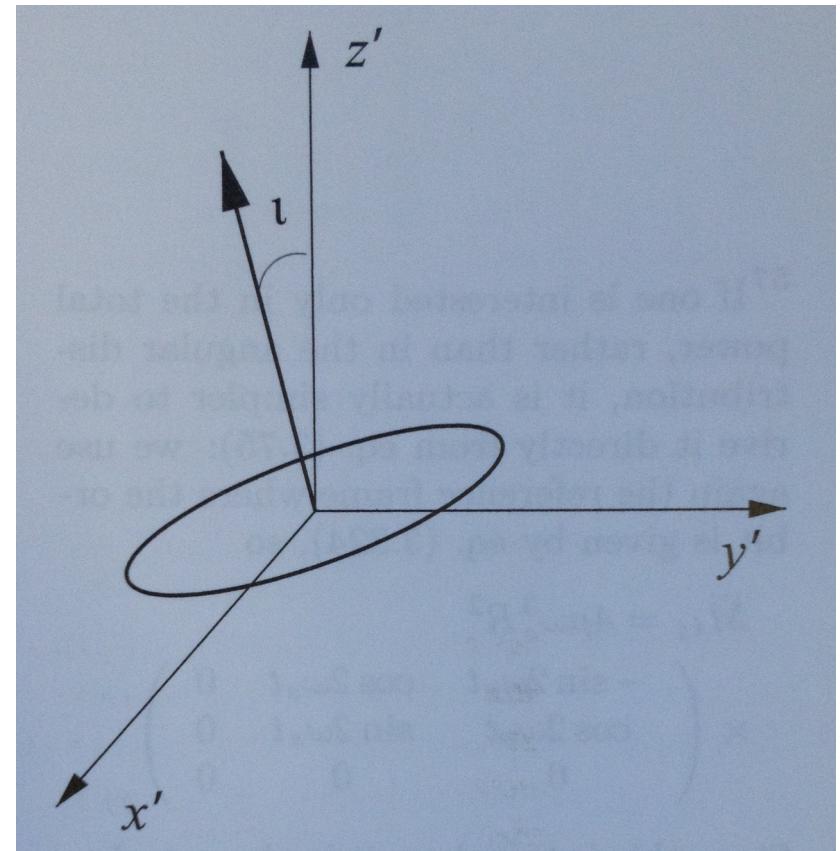


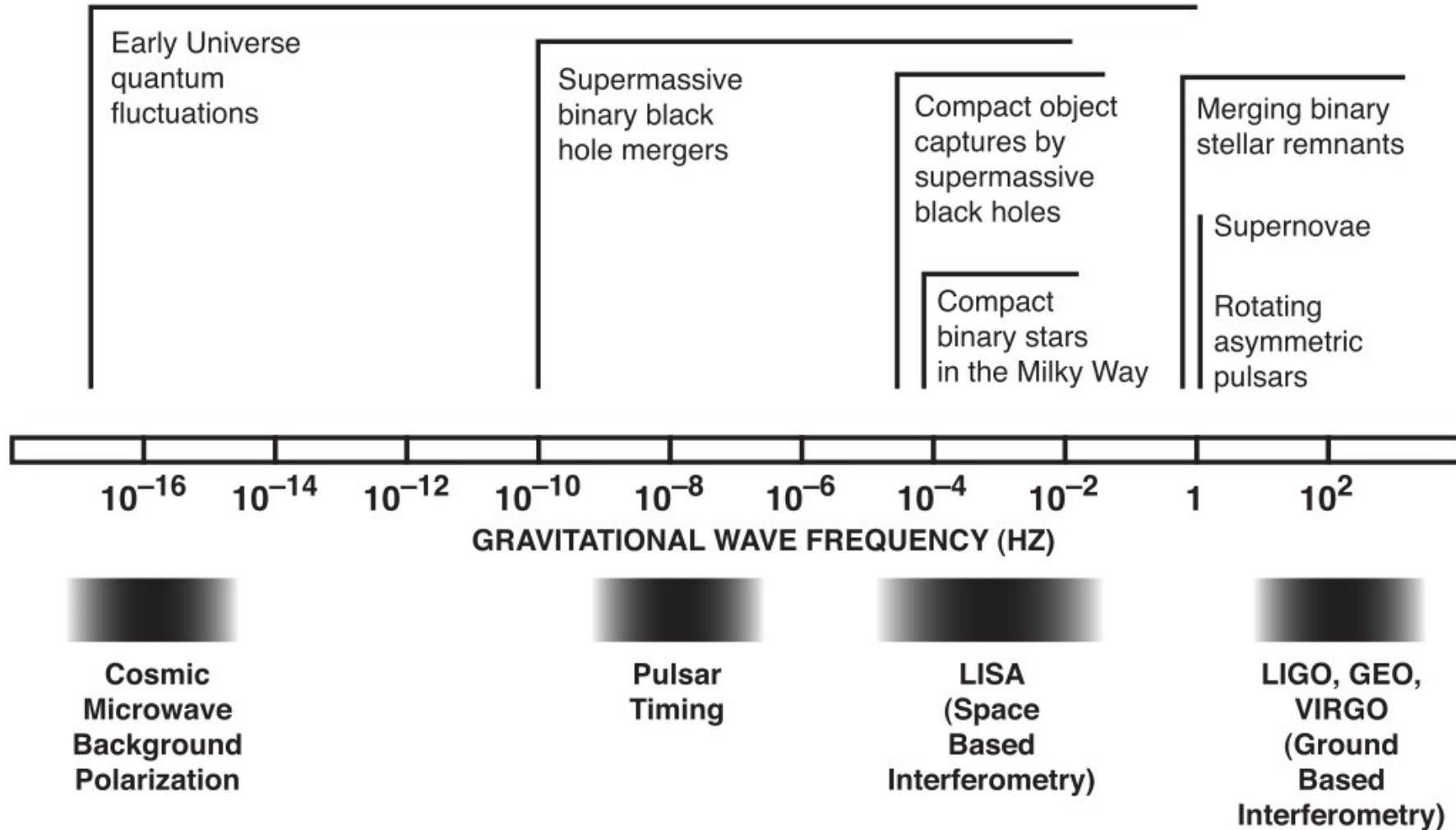
Fig. 3.6 The geometry of the problem in a frame (x', y', z') where a fixed observer is at large distance along the positive z' axis. The normal to the orbit makes an angle i with the z' axis.



Some Constants

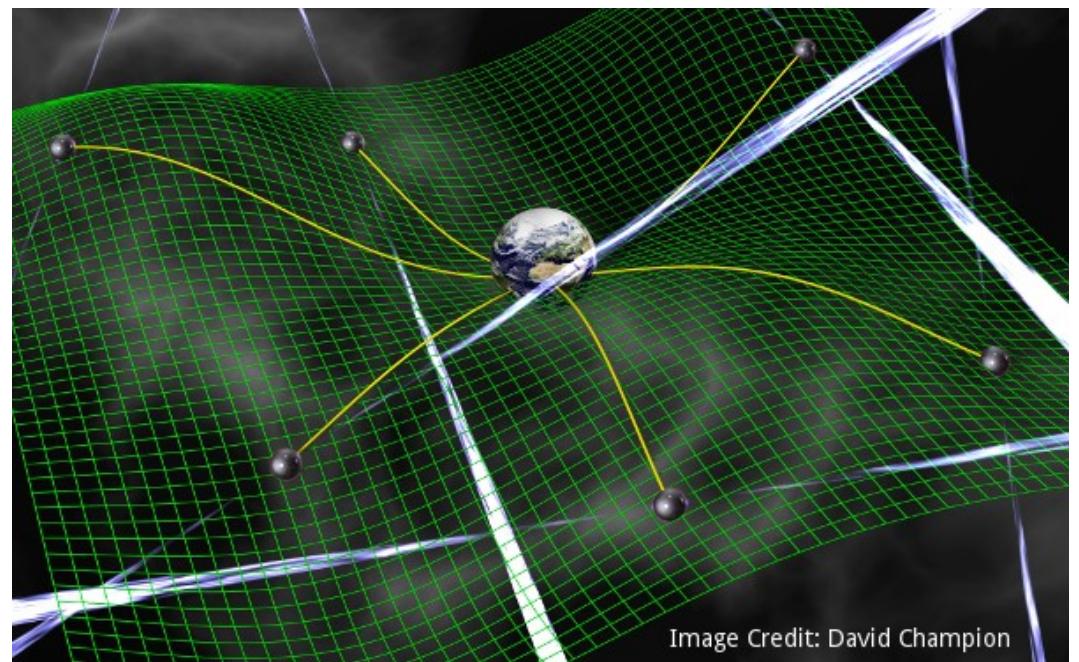
Schwarzschild radius	$2GM/c^2$	2953 m for Msun
GR Units, Mass	GM/c^2	1477 m for Msun
GR Units, Power	c^5/G	3.628e52 W
GR Units, Energy	c^4/G	1.210e44 J/m
parsec, pc		3.09e16 m = 3.262 ly
astronomical unit, au		149.6e9 m
light-year, ly		0.946e16 m = 0.307 pc
fine structure constant	$e^2/(4\pi \epsilon_0)/\hbar c$	0.00730 = 1/137.04
wavelength-energy	$\hbar c$	1.24 eV μm

Gravitational Wave Spectrum



Pulsar Timing Arrays

- Millisecond pulsars are very good clocks scattered around the Universe.
- By measuring them once in a while, you can detect a hiccup in the reception of the pulse at Earth. With enough “clocks” you can decide if there was a gravitational wave and from where it may have come.
- See chart, sensitive to very, very long wavelengths, and thus very, very small frequencies. Size of the Universe sort of waves!



Ref: [NANOGrav](#) collaboration

Pulsar Timing Arrays

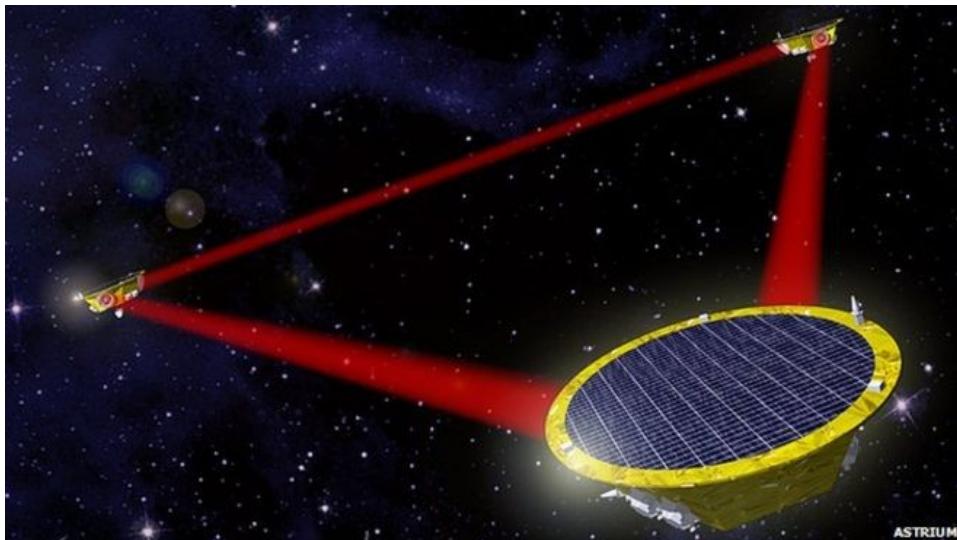
- Several collaborations working at radio astronomy facilities to measure millisecond pulsars:
- [NANOGrav](#), US and NASA, using ~~Arecibo~~ and Green Bank.
- The [International Pulsar Timing Array](#) uses many radio telescopes.
- The [European Pulsar Timing Array](#)
- The [Parkes Pulsar Timing Array](#) (part of the IPTA)
- and others.

Space-Based

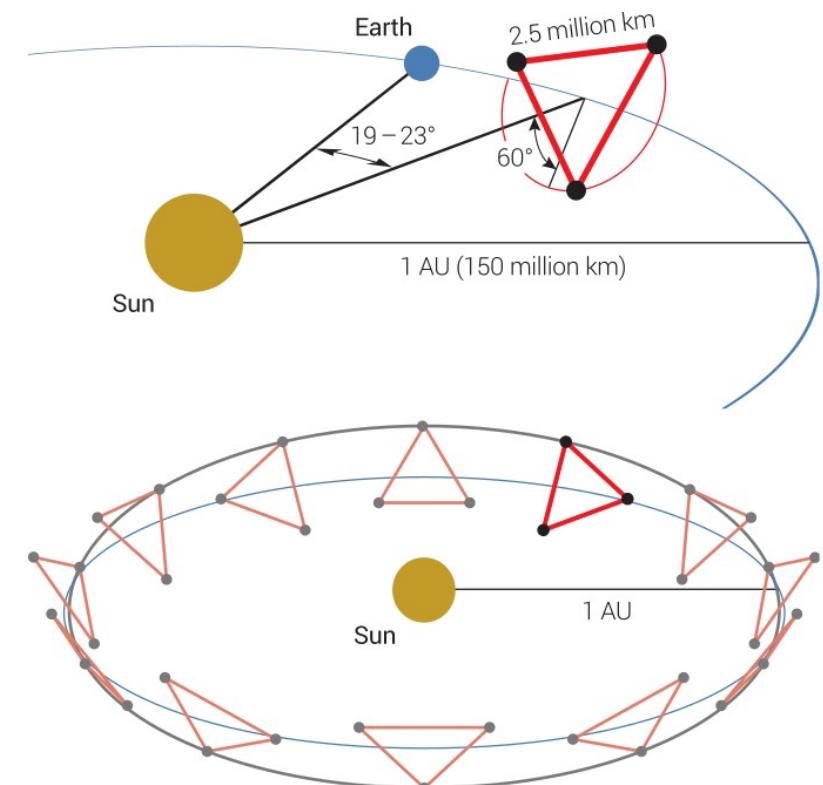
- **Laser Interferometer Space Antenna, LISA**, is a set of satellites with lasers bouncing in a triangle arrangement.
 - No seismic effects!
 - Very long baseline, 2.5×10^9 m each leg.
 - Control of craft subtle, “drag-free flying” with free test masses inside as reference.
- Currently mostly European LISA Mission with NASA/US support. LISA could launch in 2032.
 - In the hardware testing/development and design phase, soon to choose the aerospace company that will build the spacecraft.
- LISA Pathfinder satellite experiment tested several components for eLISA, early 2016.

LISA

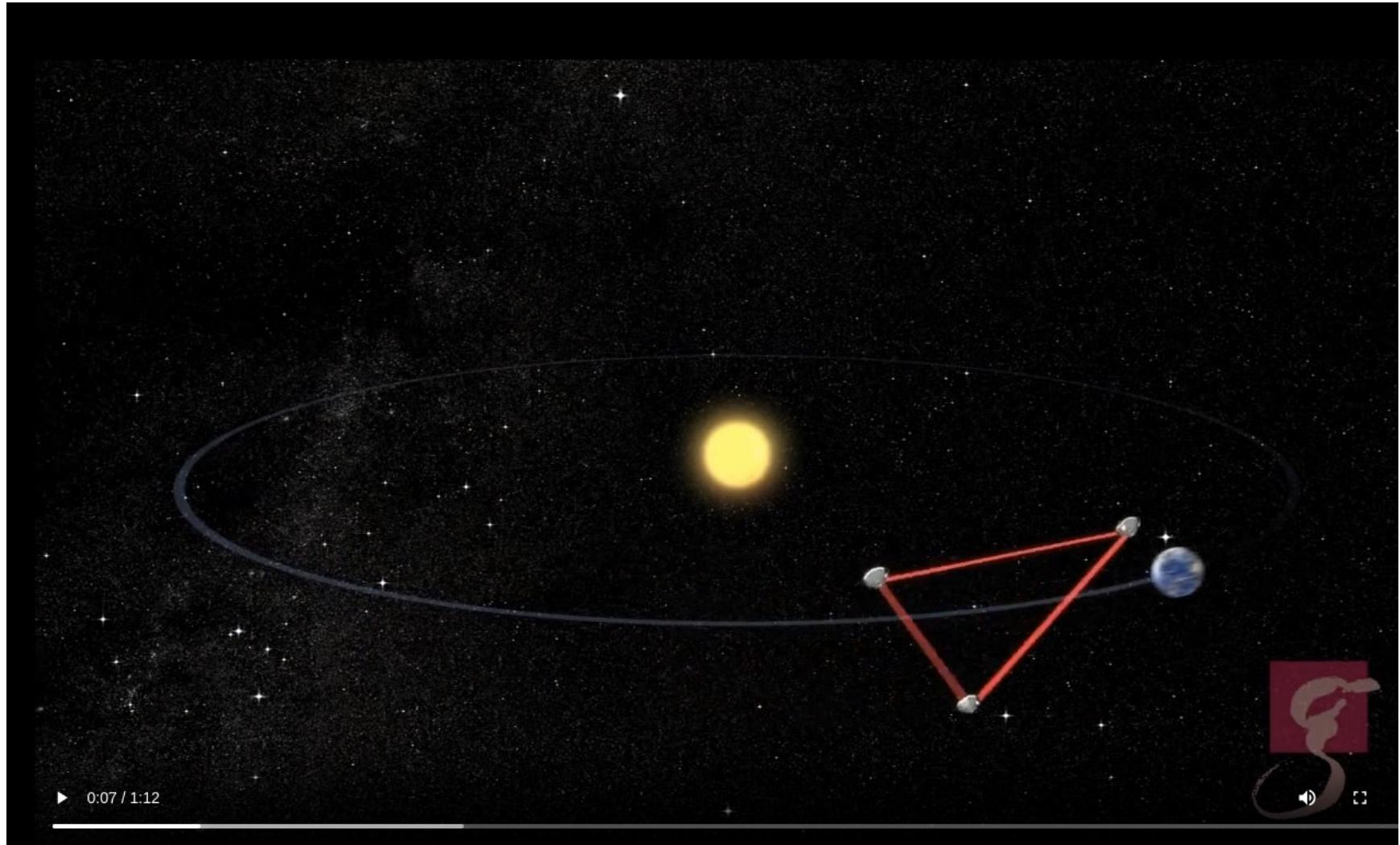
- LISA has a full equilateral triangle of lasers and baseline of 2.5e9 m. ESA mission with NASA participation.



- Chinese have two groups working on this idea:
 - TianQin
 - TAIJI

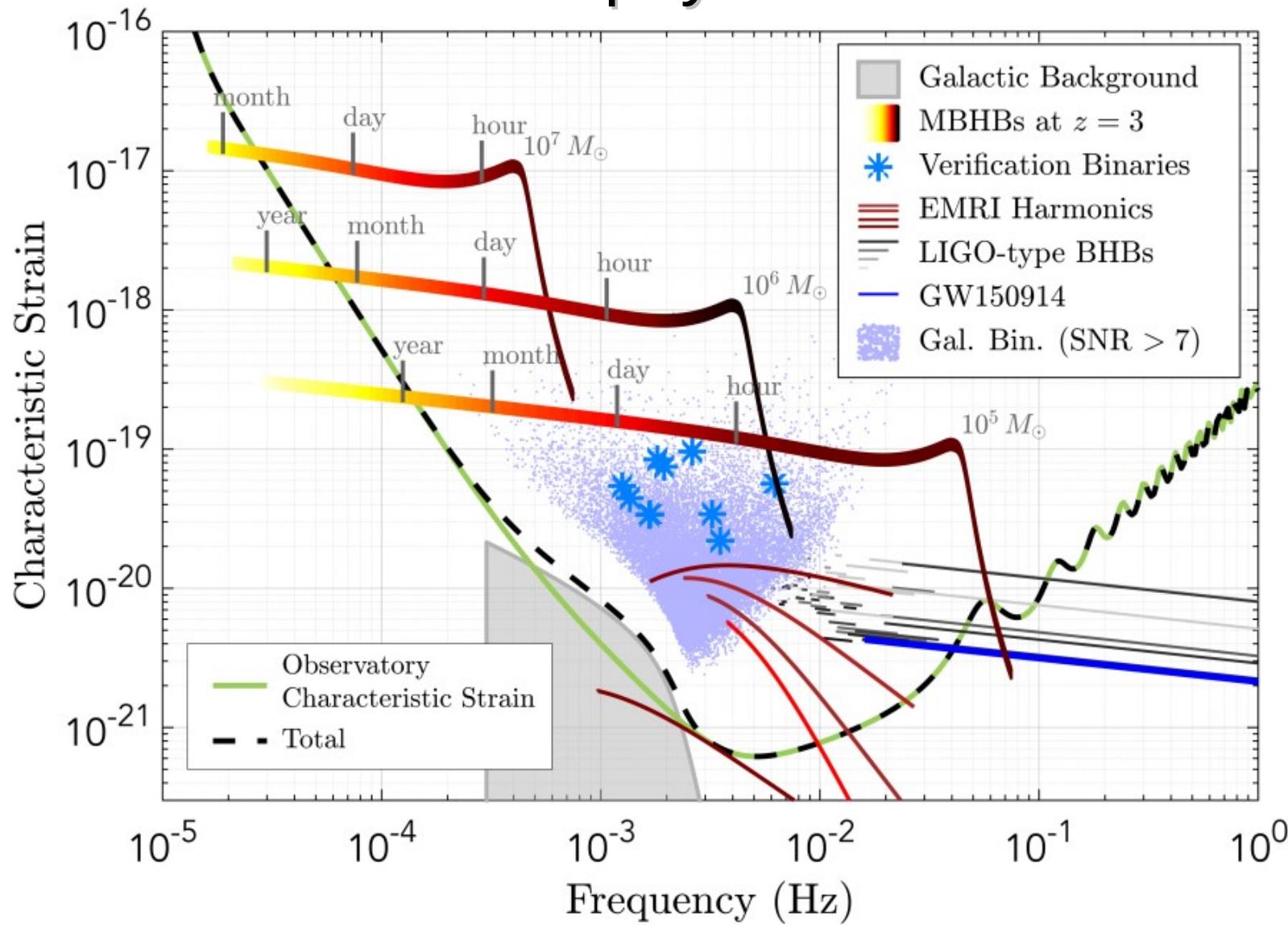


LISA Orbit



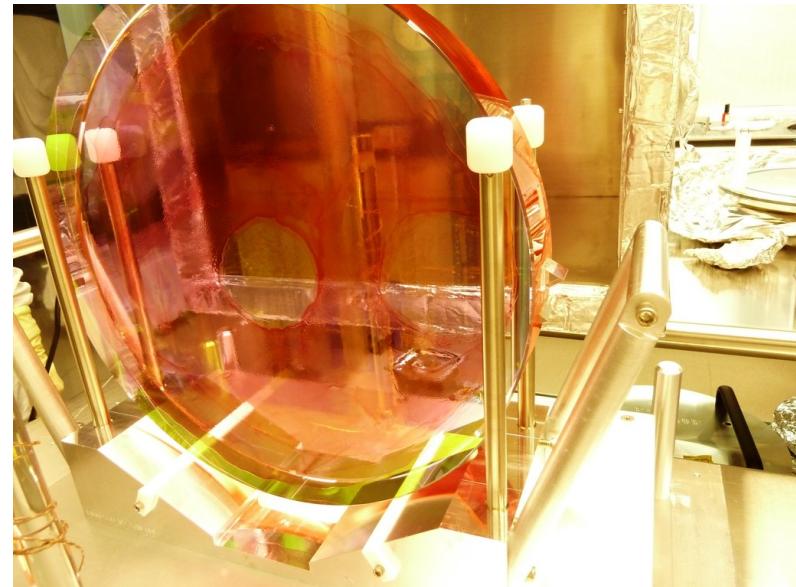
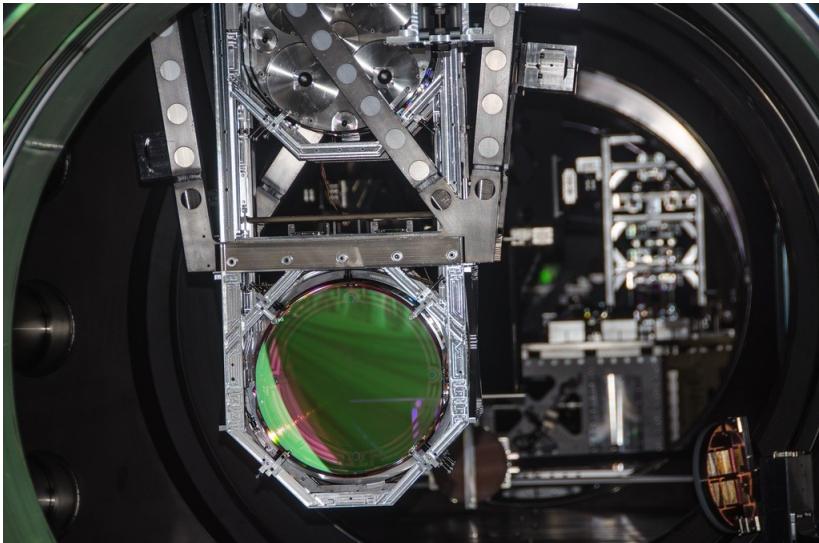
https://svs.gsfc.nasa.gov/vis/a030000/a030500/a030569/slide_09-A1_LISA_orbits2.mp4

LISA Astrophysical Sources



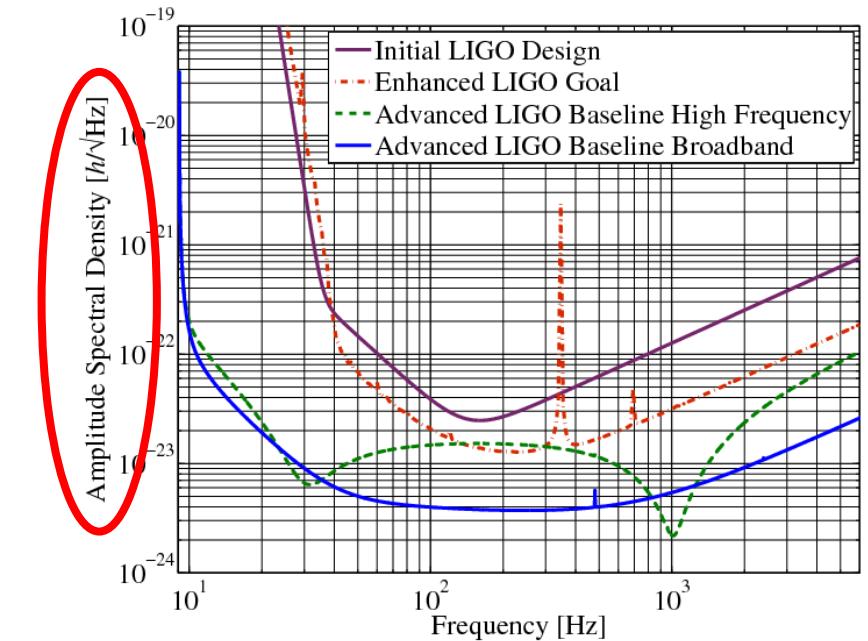
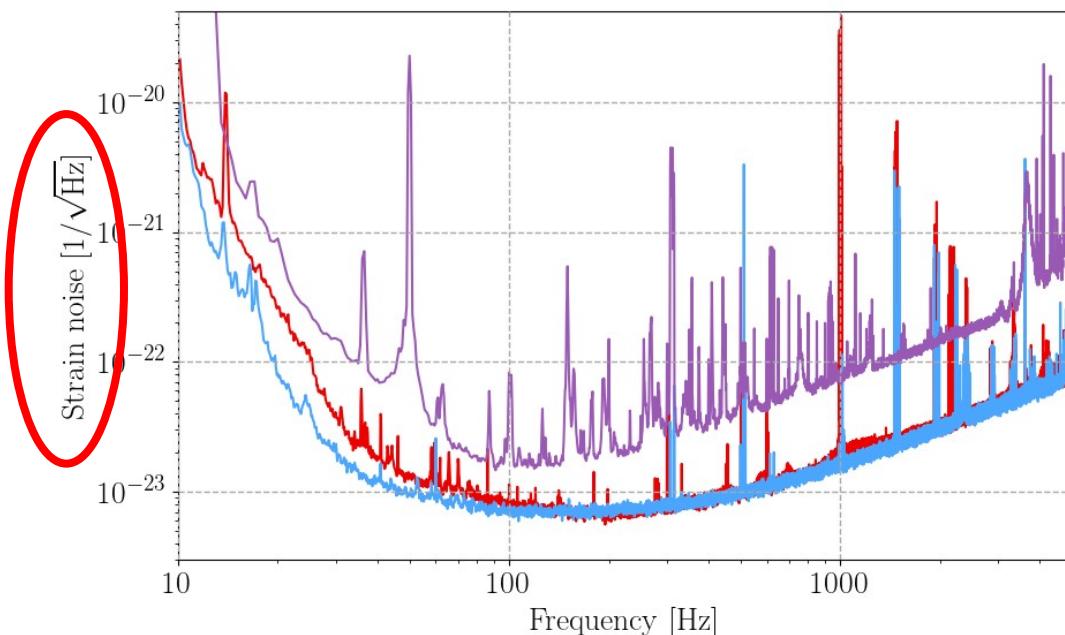
LIGO Hardware

- LIGO “Test Mass” in suspension



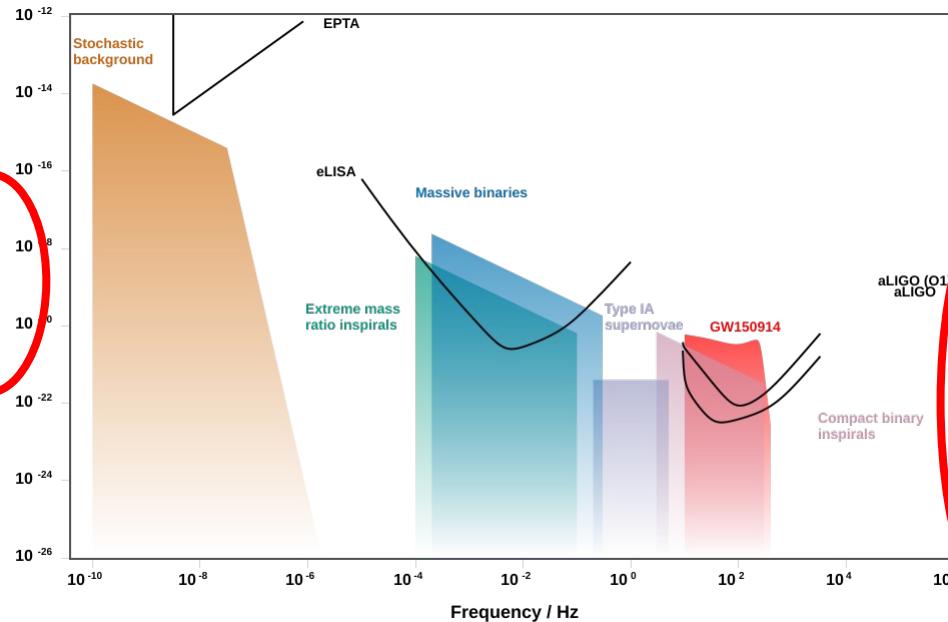
- Beam Splitter

Strain Curve, Strain Sensitivity, Amplitude Spectral Density, Power Spectral Density...OH MY!

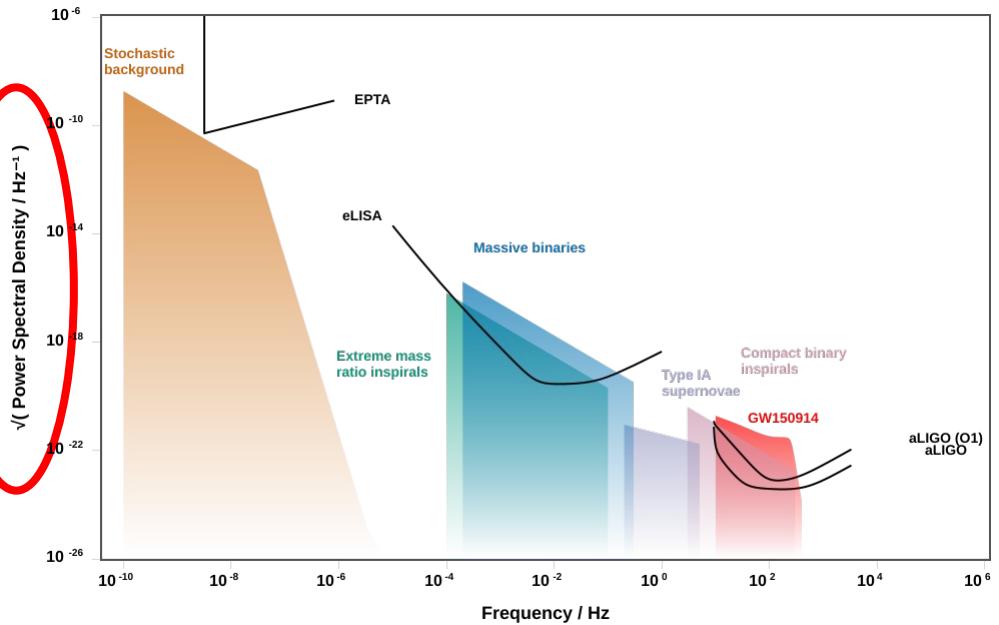


“per root Hz” is usually called the *strain sensitivity*, or sometimes *amplitude spectral density*

Strain Curve, etc...



Characteristic Strain



square root of *Power Spectral Density* is the
amplitude spectral density

Ref: Cole's online [app](#).