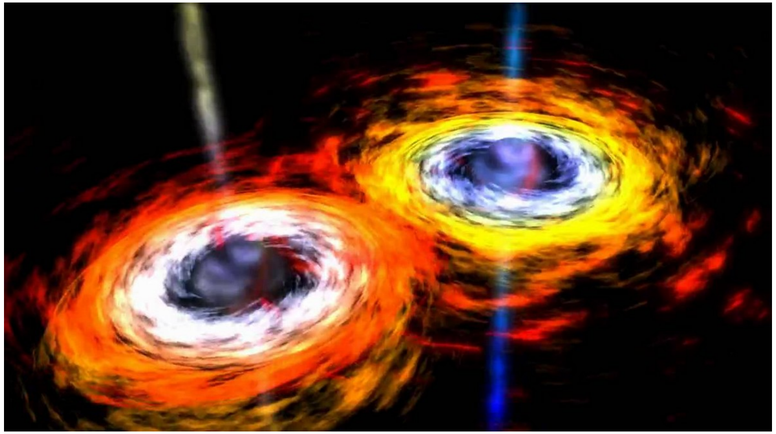
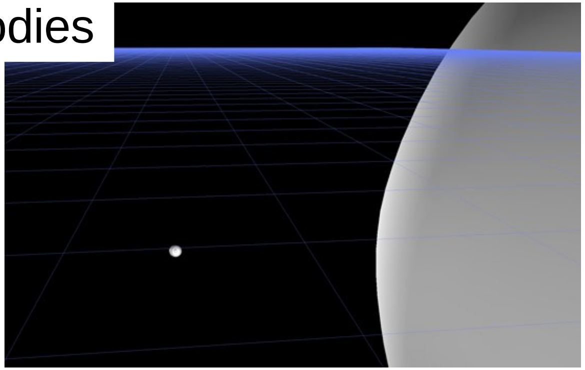
Gravitational Waves – Sources and Detectors



Orbiting Massive Bodies

Supernova, NGC7293-Helix Nebula



Extreme Mass Ratio Black hole Mergers

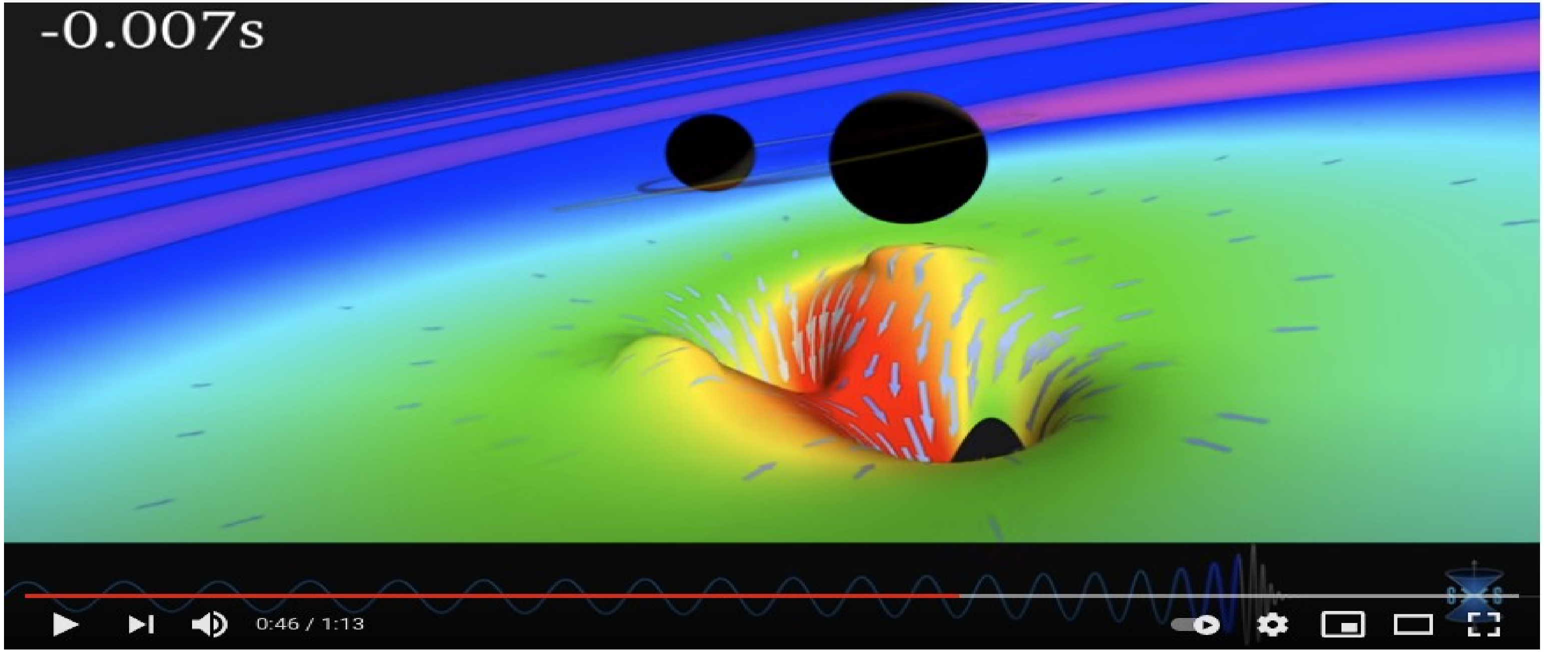
Credit: https://www.rit.edu/news/story.php?id=47936

Sources of Gravitational Waves

- **Ripples in space time**...hard to do, spacetime is very “rigid” or “stiff.”
- Need **matter in motion** to create Gravitational Waves, recall charges in motion create electromagnetic waves.
- The motion **cannot be spherically symmetric**. Conservation of mass and conservation of linear momentum (of the mass) means the first “moment of the mass distribution” that can cause GWs is “quadrupole.”
- **Large masses and fast changes in time** create stronger GWs.
  - Black holes are small and massive. In binary orbits with another black hole, orbital speeds can reach more than 30% the speed-of-light! Small is good because they can get very close before they “touch” or “merge.”
- **Soooooo....**

Sources, Binary Black Holes

- **Binary Black Holes** inspiral, merger, and ringdown. First discovered/seen **GW150914**, announced April 2016.



https://youtu.be/1agm33iEAuo 1m13s, no matter around, GW150914 parameters

Sources, Binary Neutron Stars

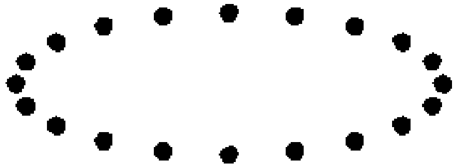
- **Binary Neutron Stars** also inspiral, merger, and ringdown. First **GW170817** . Seen in all optical and other EM bands.

https://youtu.be/bBCArmUPgCw 48s, matter around, kilonova, heavy elements created.

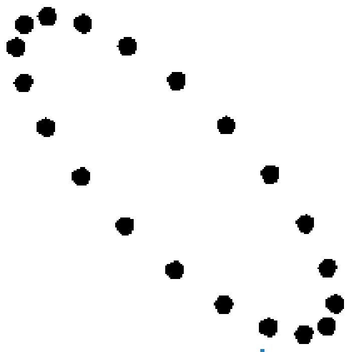
Nature of Gravitational Waves, and $h_{\mu\nu}$

Of all 10 possible components to h, it simplifies to two polarizations. One with a motion that shrinks (grows) in one direction while growing (shrinking) in the other direction. The other polarization is just rotated by 45 degrees.

h_+



$h_\times$



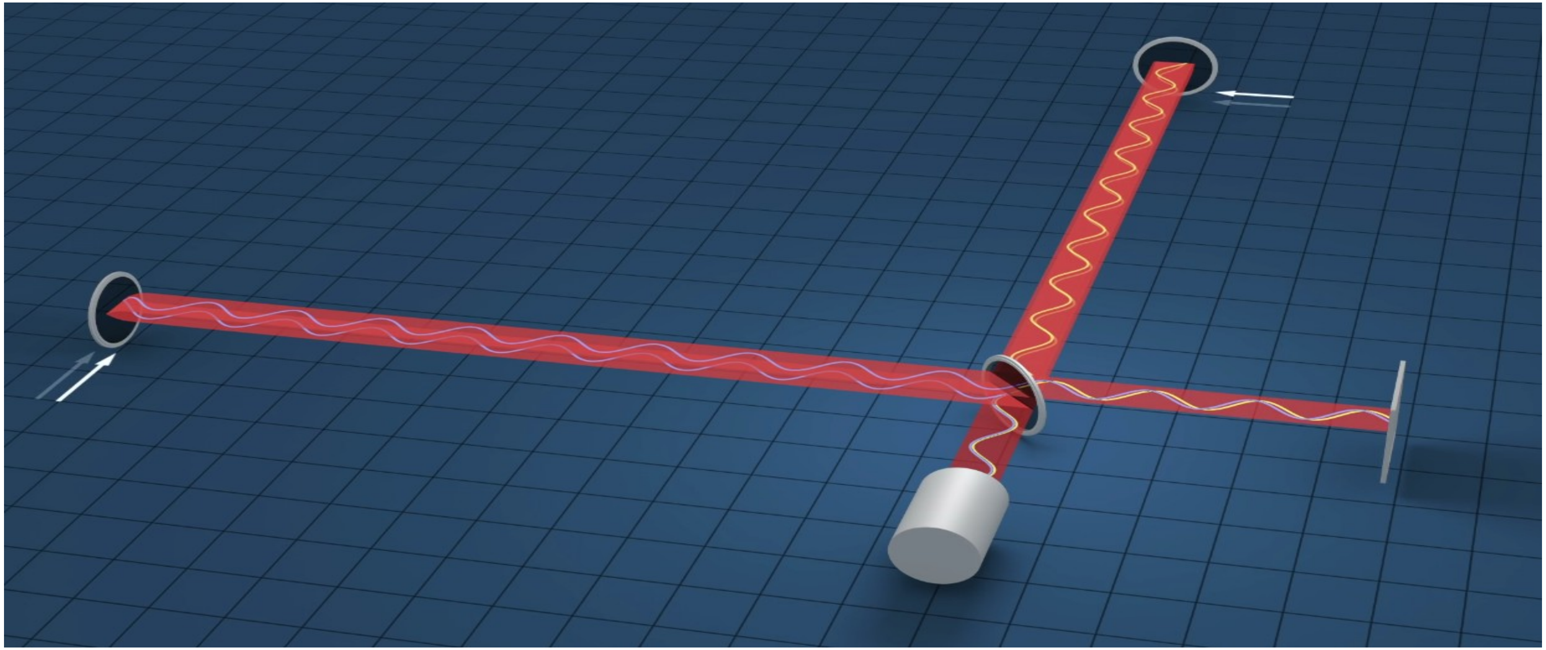
hplus polarization

hcross polarization

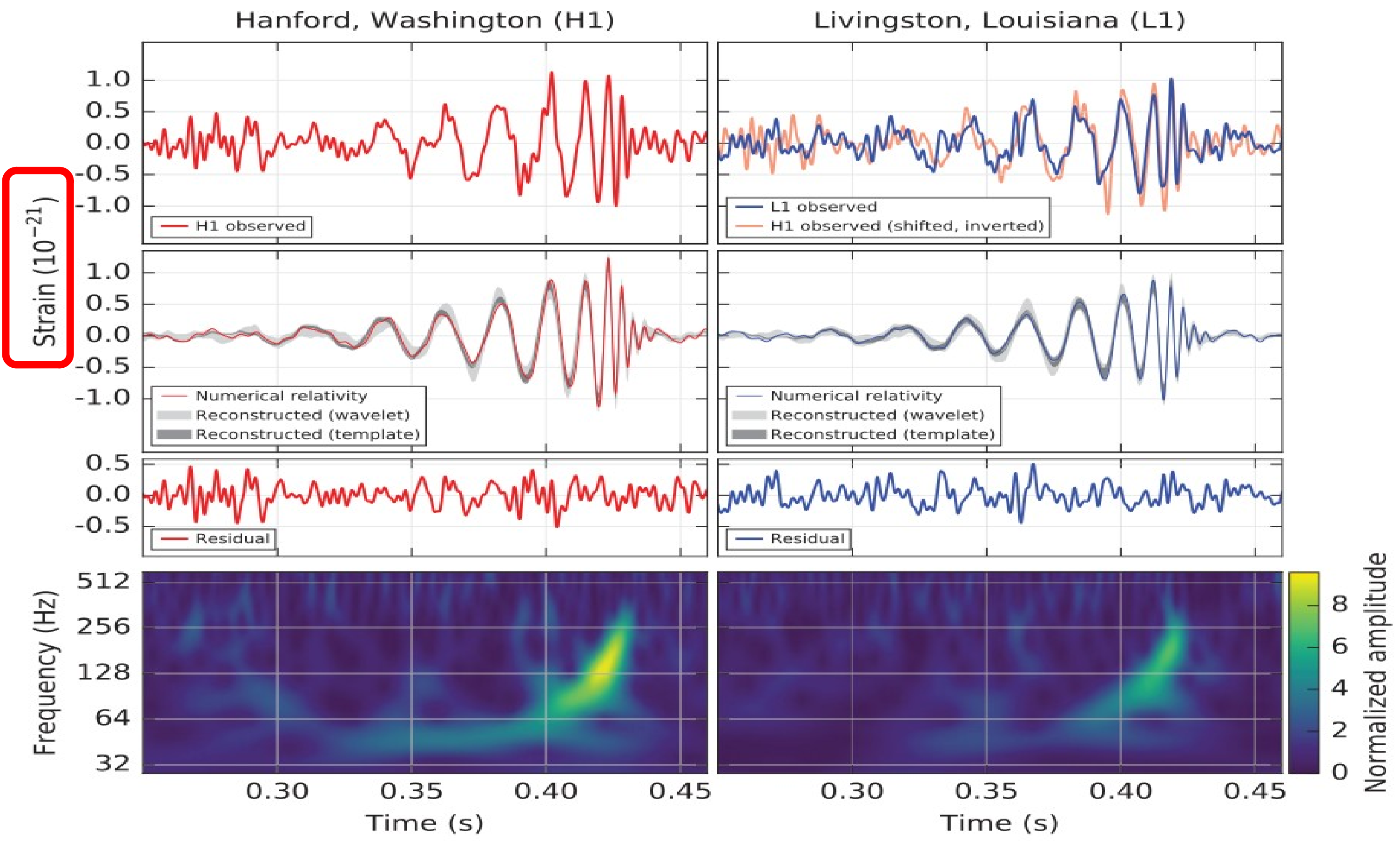
Wave headed into the page, Y up, X left.

Big Michelson Interferometer, 4 km arms

- YouTube link https://youtu.be/tQ_teIUb3tE



Gravitational Wave GW150914



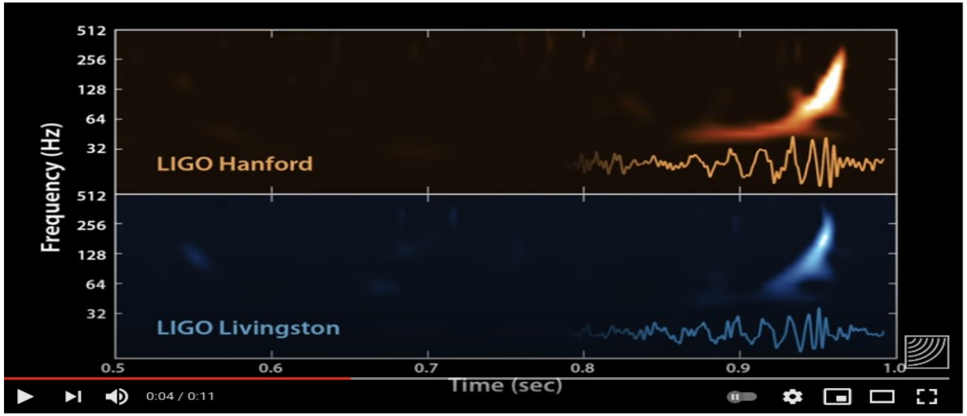
Direct Discovery by LIGO of a gravitational wave announced on 11 February 2016 for a wave on 14 September 2015---event labeled GW150914.

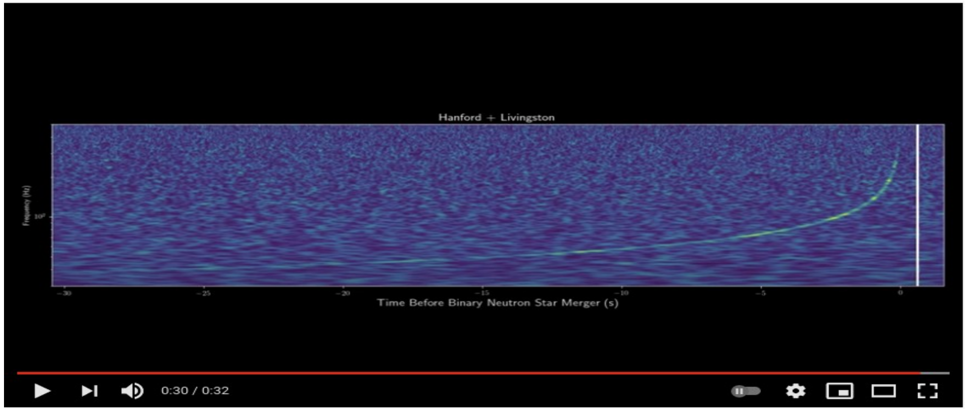
SNR = 24, 5 sigma detection!

Pretty after filtering out the known seismic frequencies.

Sources, Waveforms---The Strain

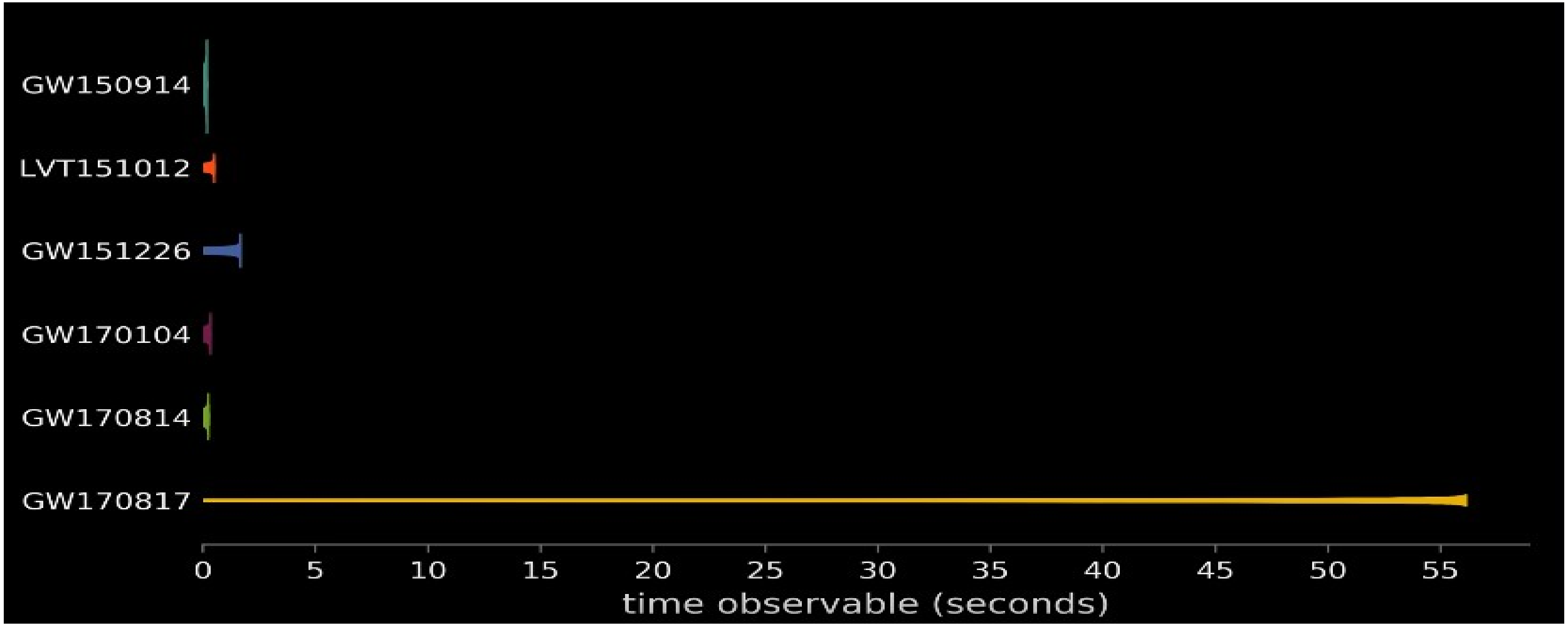
- **Waveforms**, the strain in the interferometer(s)
- Binary Black Holes, GW150914, here





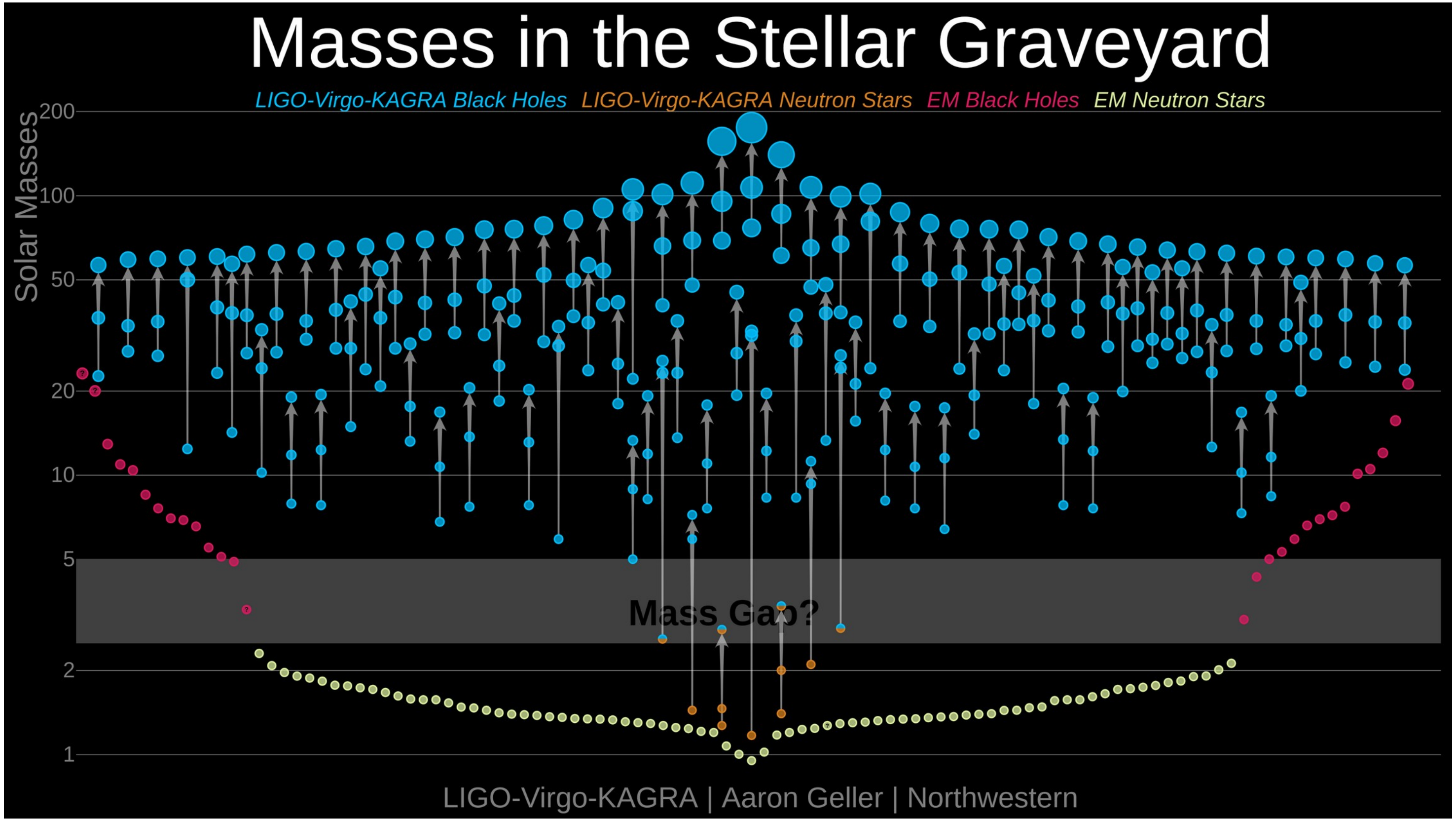
- Binary Neutron Stars, GW170817, here

GW170817 with other BH GW Sources



YouTube video, animated waveforms, here

Population of Black Holes and Neutron Stars



https://media.ligo.northwestern.edu/

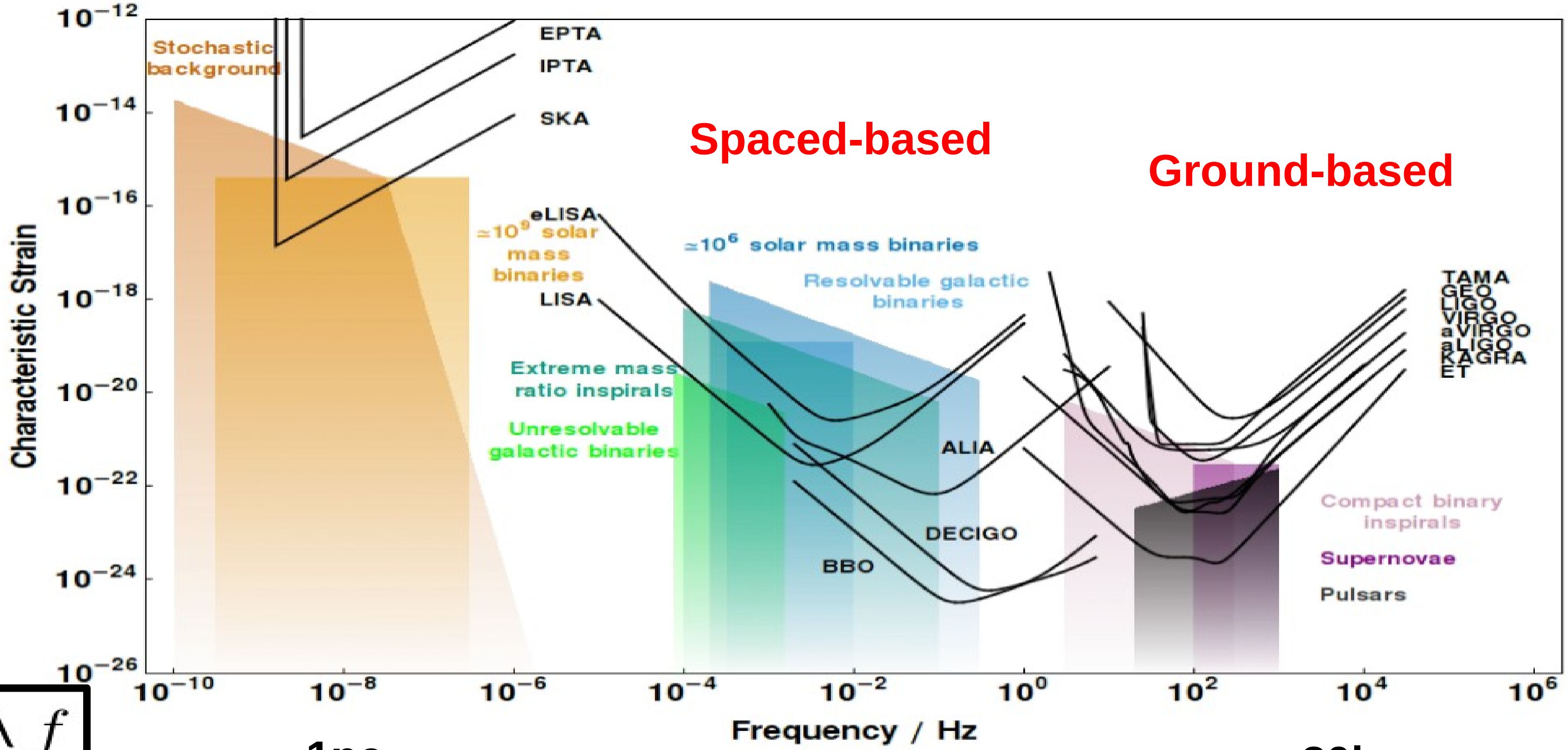
Reference to 2020 Image

above in 2022

Strain Curves from Moore et al.

Pulsar Timing Arrays

http://gwplotter.com/



$$\boxed{c = \lambda f}$$

wavelength, m	**1pc** **3e16**	**3e12**	**3e8**	**30km** **3e4**

LIGO, ground-based Gravitational Wave Detectors

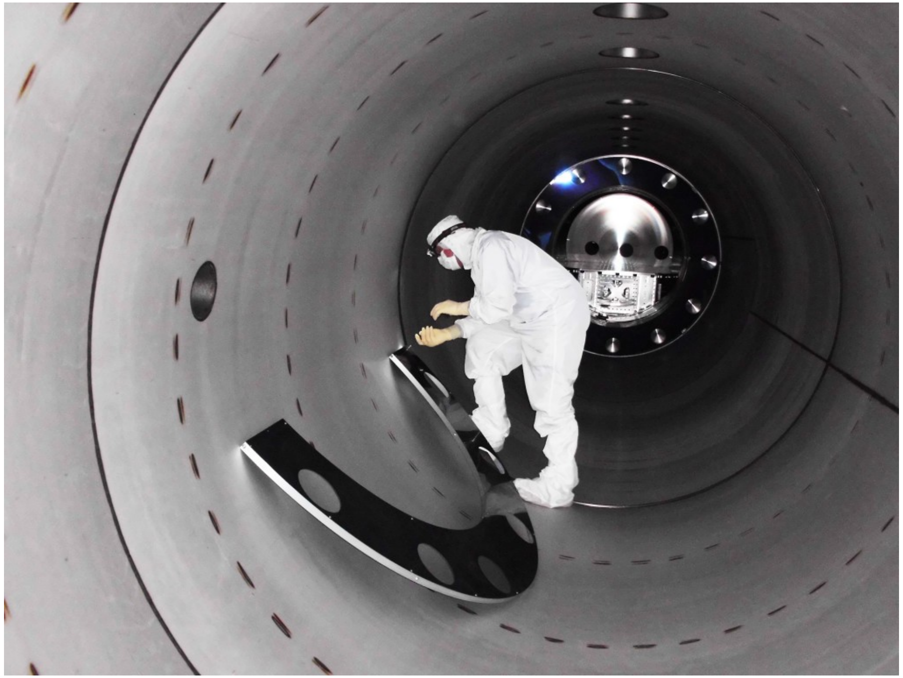
LIGO Livingston, LA (LLO)

LIGO – Laser Interferometric Gravitational Wave Observatory

LVK – LIGO, Virgo (Italy), KAGRA (Japan) collaboration, sometimes LSC

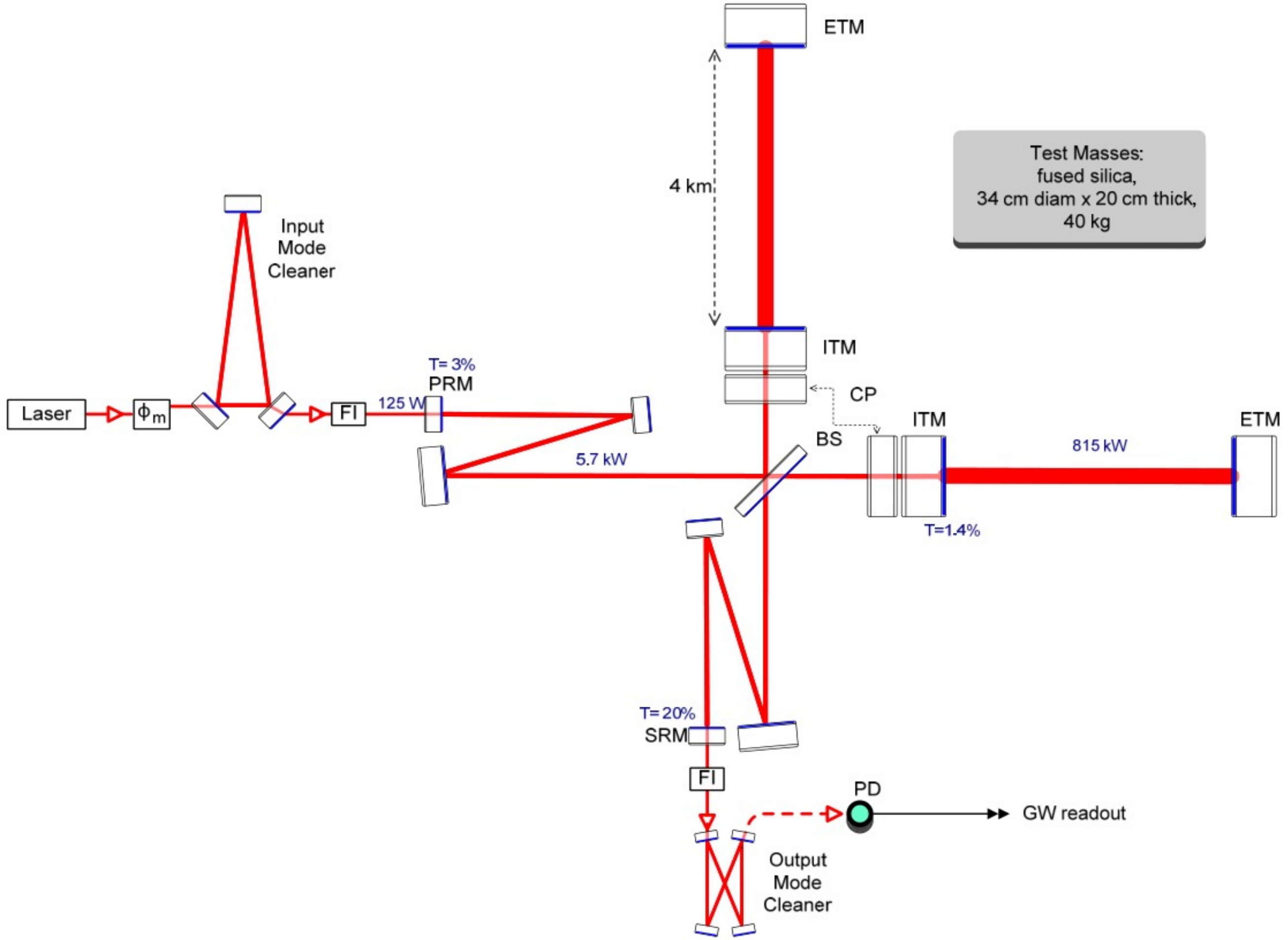
LIGO Hardware

- LIGO “Test Mass” in 4-element suspension



- LIGO Vacuum Tube

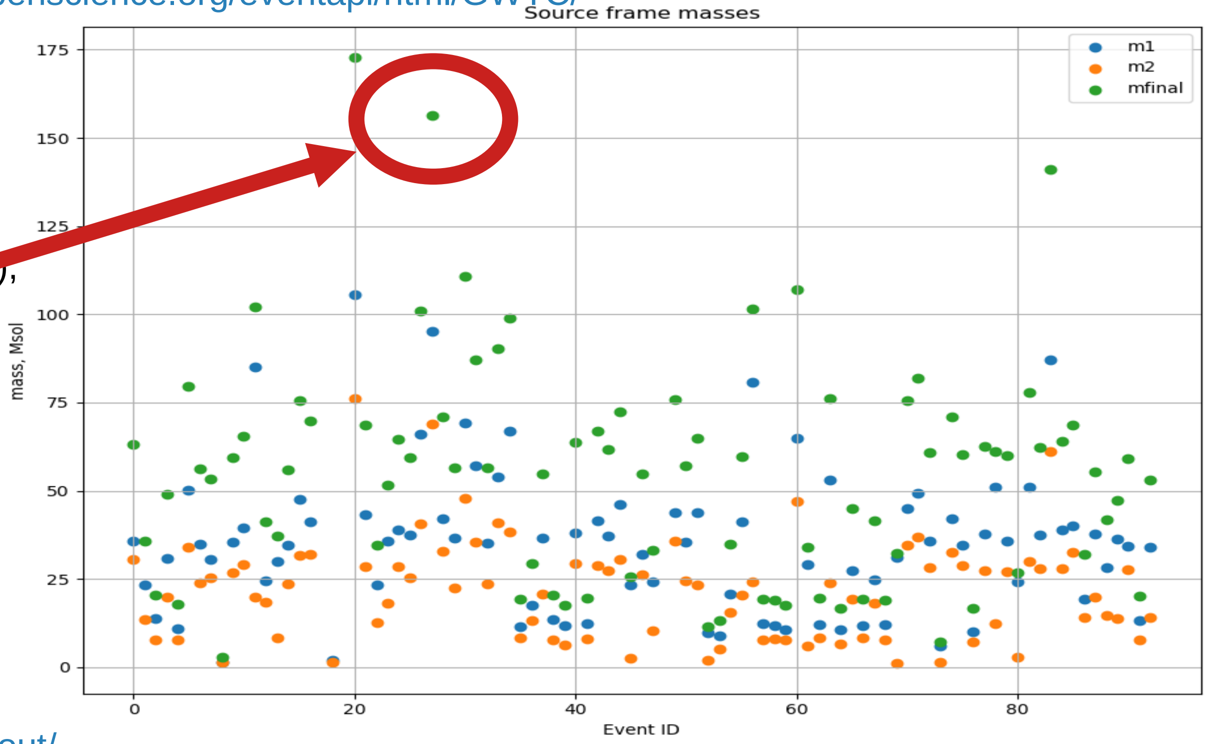
Ref: some hardware images


ETM
4 km
Test Masses:
fused silica,
34 cm diam x 20 cm thick,
40 kg
Input
Mode
Cleaner
ITM
T= 3%
PRM
CP
Laser
ϕ_m
FI
125 W
ITM
ETM
BS
815 kW
5.7 kW
T=1.4%
T= 20%
SRM
FI
PD
GW readout
Output
Mode
Cleaner

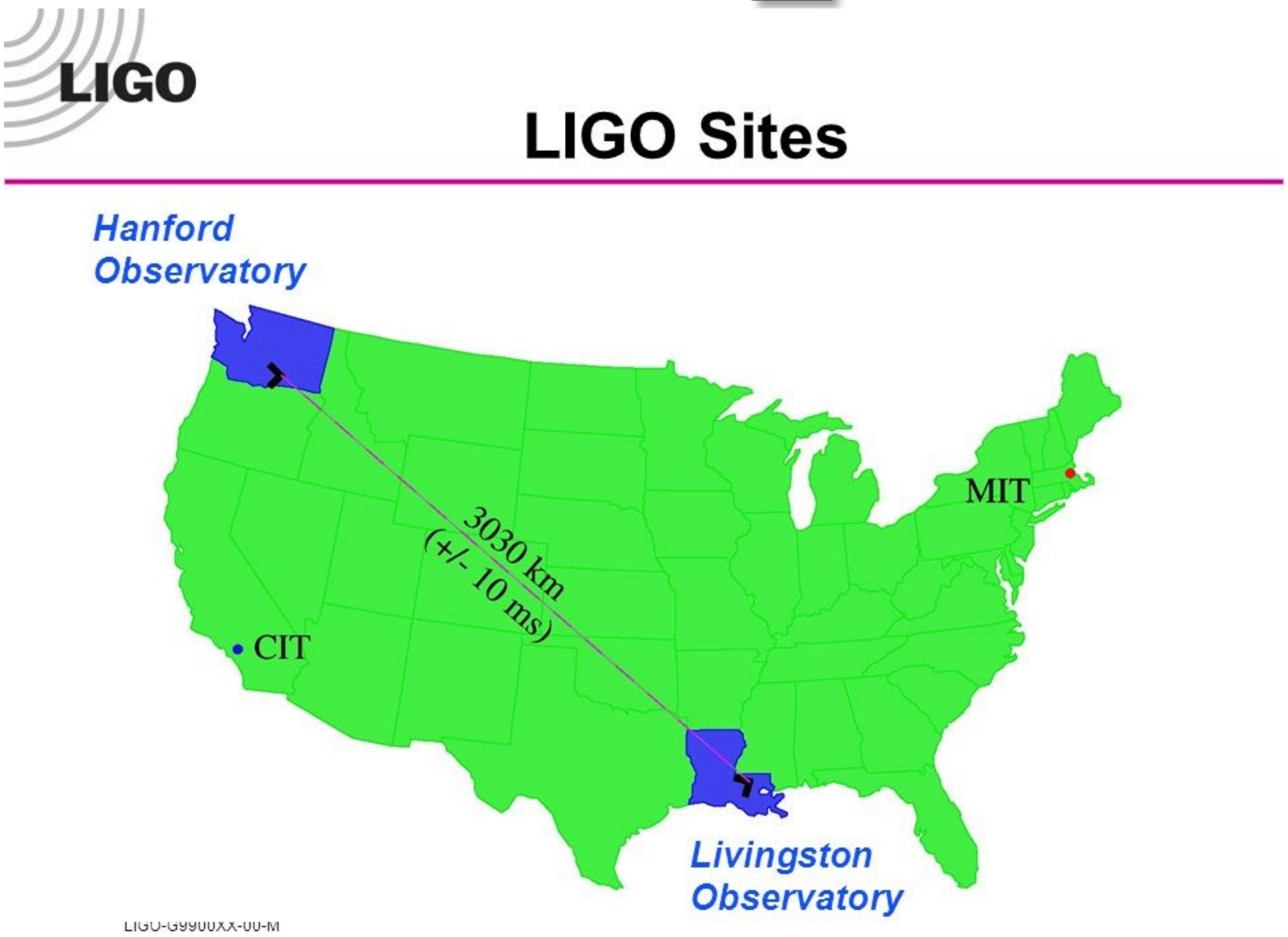
Catalog GWTC-3

The publicly released table of gravitational wave events, see the page
https://www.gw-openscience.org/eventapi/html/GWTC/

Event **GW190521_v3**, first Intermediate Mass Black Hole (IMBH), 156.3 Msol

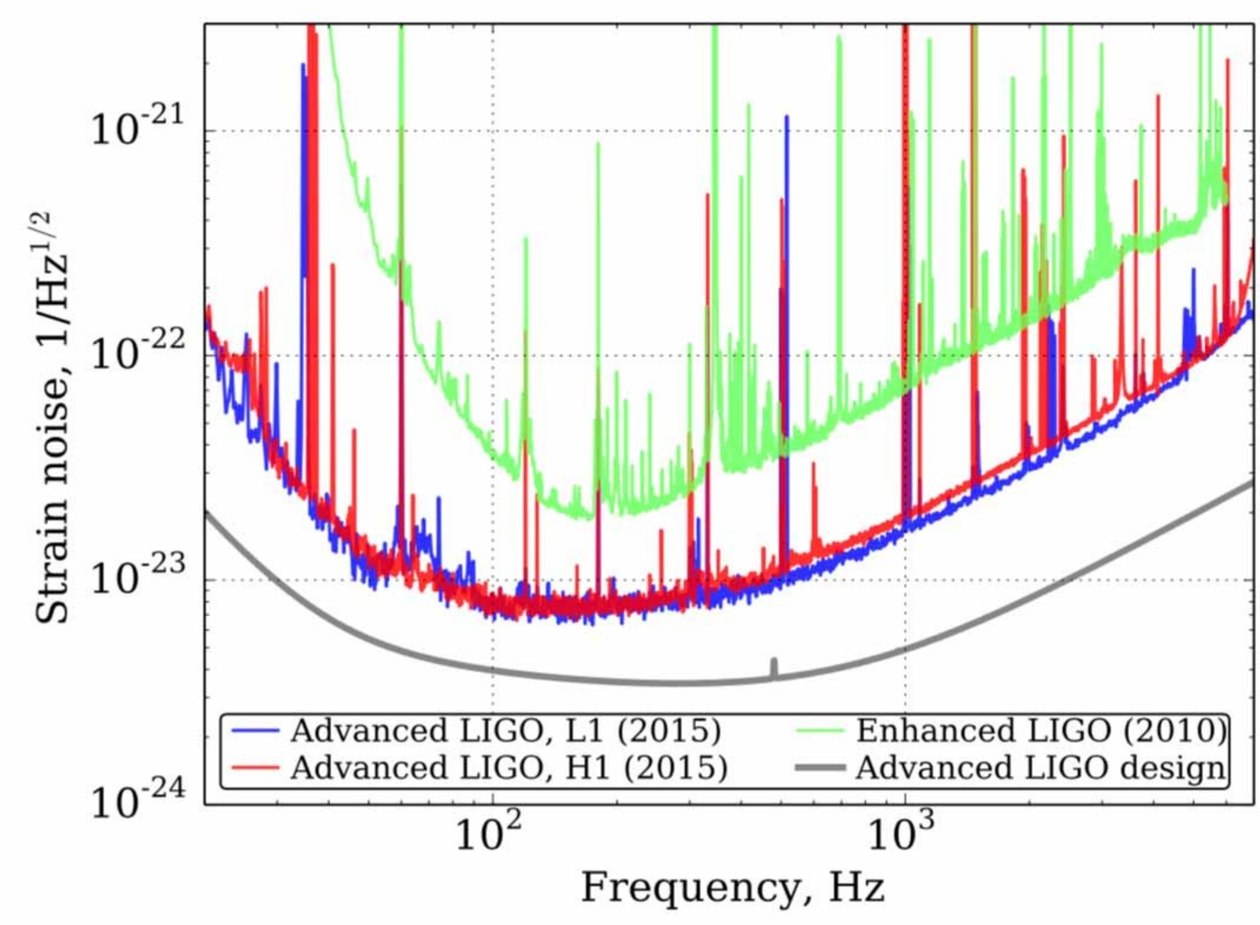


https://www.gw-openscience.org/about/



Also Virgo, Cascina, Italy and
KAGRA, Japan

Seismic Noise aLIGO and initial



Above the background noise.

Below the noise.

Summary

- Vanderbilt University (Kelly Holley-Bockelmann, Karan Jani, and group) are in LIGO, since September 2020.
  - Led by Karan Jani, newly hired assistant professor at Vanderbilt, researching Intermediate Mass Black Holes and Lunar based GW detectors.
- LIGO and KAGRA started taking data again May 2023.
  - Called O4.
  - Virgo (Europe) is finishing upgrades, and will join O4.

Thank you for your Attention!

Questions? Comments?

Links

- LIGO Teacher's page here
- LIGO Student's page here
- LIGO CalTech Educational Resources here
- LIGO Scientific Collaboration, LSC
- LIGO Document Server
- CalTech GW media assets page.
- Kelly Holley-Bockelmann's TEDx Nashville Talk 2016 about GW150914.
- Pulsar timing array, Nanograv
- LISA Consortium, space-based GW detectors

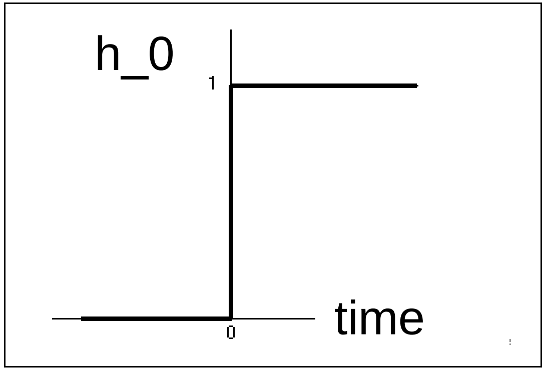
Backup

Saulson, *If Light Waves…*

- We know in the expanding Universe, light continuously experiences the increase in scale by expanding continuously itself.

$$\frac{\lambda_1}{\lambda_0} = \frac{R(t_1)}{R(t_0)}$$

- **NOT** the case for a light wave in LIGO...
  - The wave filling the cavity **IS** stretched by the sudden onset of the gravitational wave.
  - Light travel time, 2x4km is 27μs, but GW oscillates at 20 Hz, 50ms.



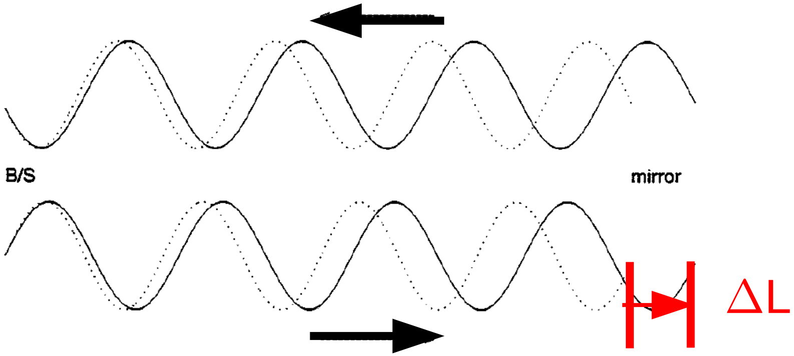


Fig. 2. Light before (dotted) and after (solid) the arrival of a gravitational wave. The beamsplitter is at left, end mirror at right. Outbound light is shown at the bottom, returning light at the top.

Ref: Saulson, *If Light Waves…*, AmJPhys **65**, 501 (1997)

As the Wave runs through...

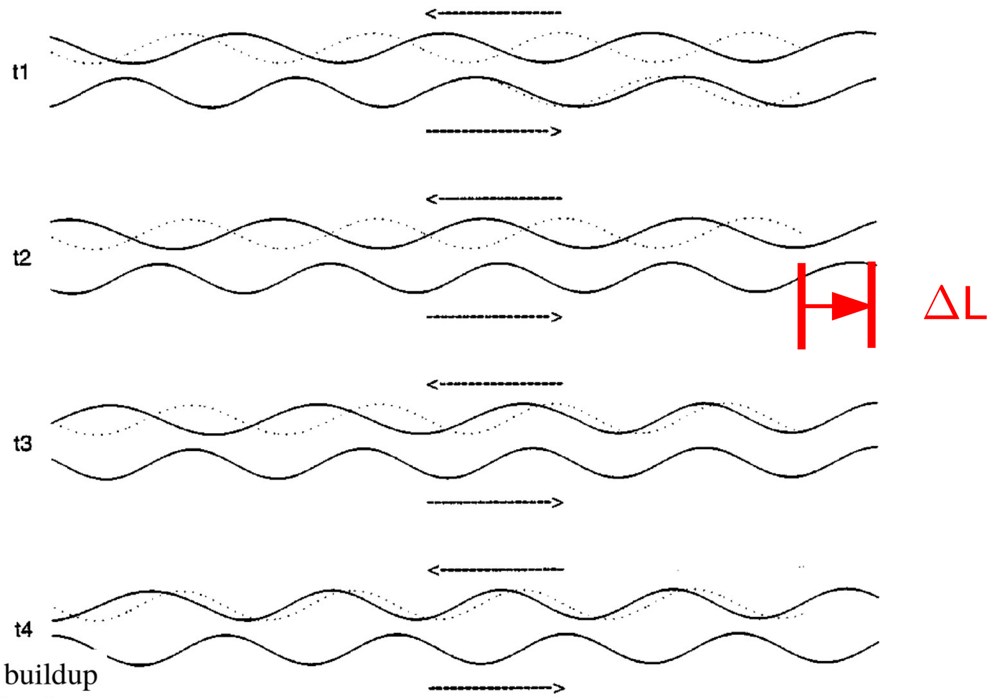


Fig. 3. Like Fig. 2, but at a succession of later moments. Note the buildup of phase shift between the light in the stretched arm (solid) compared to how it would have traveled through an unstretched arm.

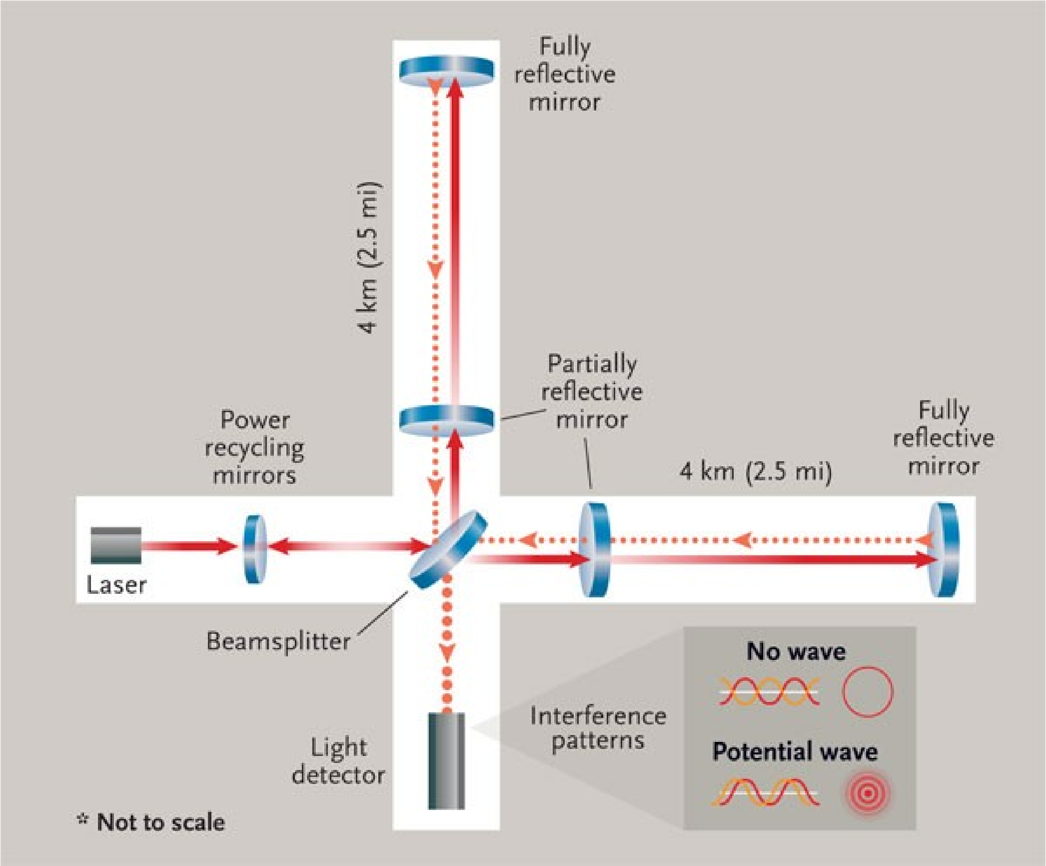
Saulson, *If Light Waves…*, says rods are still rigid!

V. LENGTHS IN COSMOLOGY AND IN LABORATORY PHYSICS

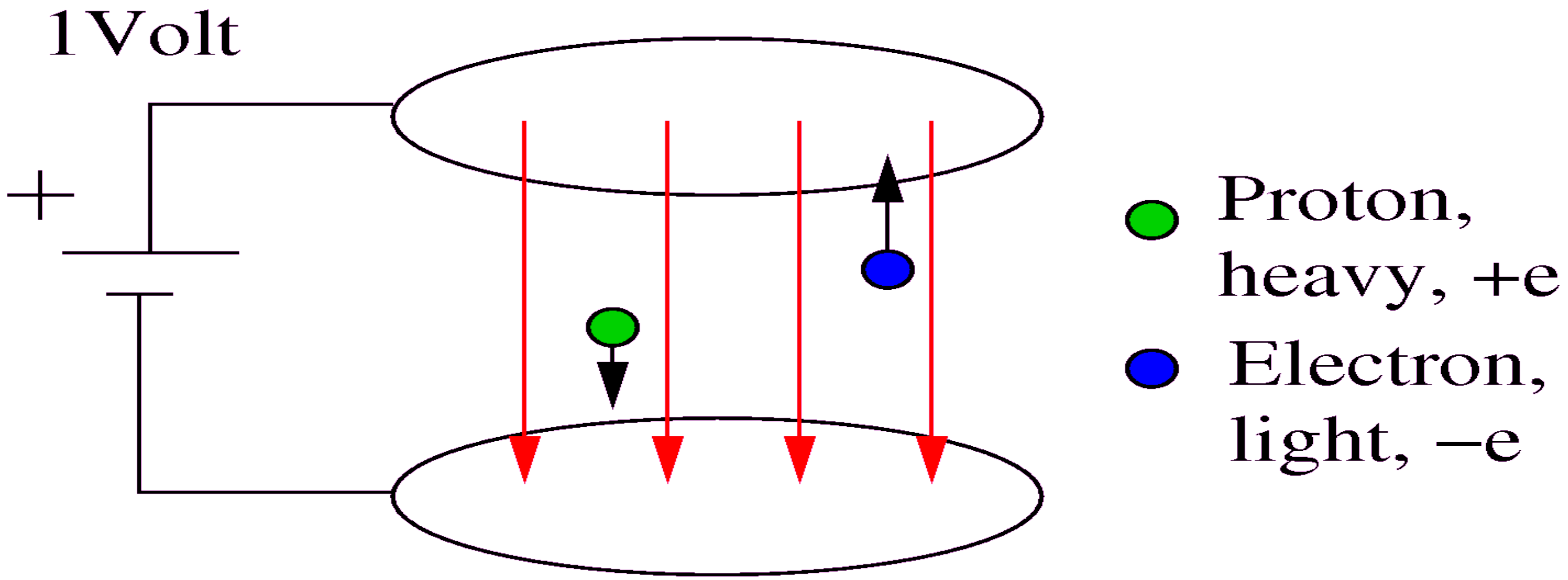
Note that the language we have been using in this paper only makes sense if we imagine that we have standards of length other than either the separations of freely falling test masses or the wavelengths of light waves. We do. A good paradigm of a length standard is a perfectly rigid rod. Such a rod does not change its length in the presence of a gravitational wave, because the arbitrarily strong elastic forces between its parts resist the gravitational force carried by the gravitational wave. As we will see below, we can also use the travel time of light as a reliable ruler under most conditions, in spite of the stretching of light waves that goes on when space expands.

LIGO Inteferometer

- 4km long



Units?



SI Prefixes

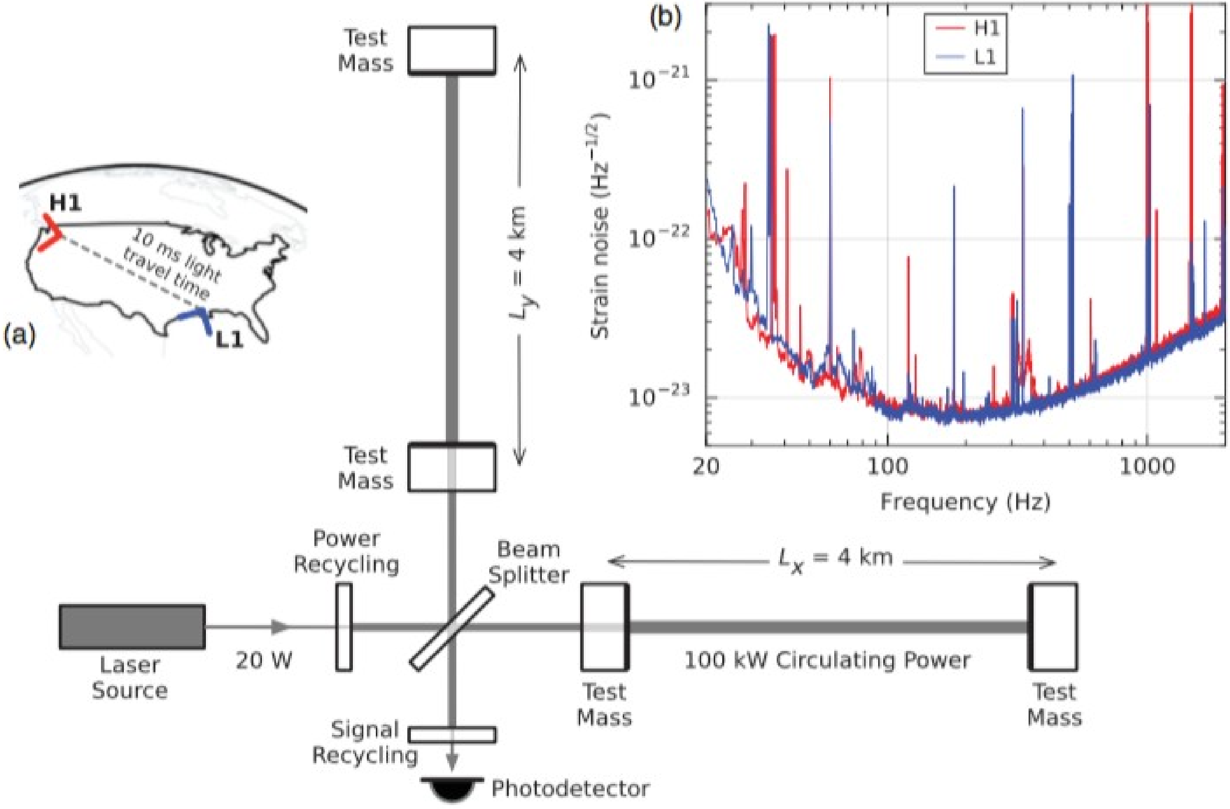
Table 5. SI prefixes

Factor	Name	Symbol	Factor	Name	Symbol
10^{24}	yotta	Y	10^{-1}	deci	d
10^{21}	zetta	Z	10^{-2}	centi	c
10^{18}	exa	E	10^{-3}	milli	m
10^{15}	peta	P	10^{-6}	micro	µ
10^{12}	tera	T	10^{-9}	nano	n
10^{9}	giga	G	10^{-12}	pico	p
10^{6}	mega	M	10^{-15}	femto	f
10^{3}	kilo	k	10^{-18}	atto	a
10^{2}	hecto	h	10^{-21}	zepto	z
10^{1}	deka	da	10^{-24}	yocto	y

Merging Black Holes

ref: http://www.techinsider.io/binary-black-holes-confirmed-gravitional-waves-2016-2

Schematic, Map, and Noise



GW strain for Circular Orbit

$$h_0 = \frac{r_{s1} \cdot r_{s2}}{r \cdot R} \qquad \frac{\omega_s^2}{c^2} = \frac{(r_{s1} + r_{s2})}{2\, R^3}$$

$$h_+(t) \quad = \quad h_o\left(\frac{1+\cos^2 i}{2}\right)\, \cos 2\omega_s t$$

$$h_\times(t) \quad = \quad h_o\, \cos i\, \sin 2\omega_s t$$

- rs1 is 2GM_1/c^2, Schwarzschild radius for mass M_1, etc.
- r is distance from Earth to system.
- R is the separation of the two bodies.

ref: Maggiore around Eqn. 3.332.

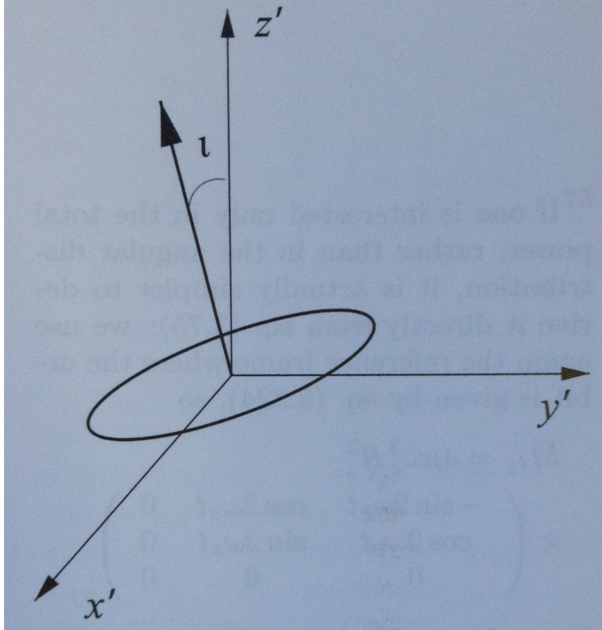
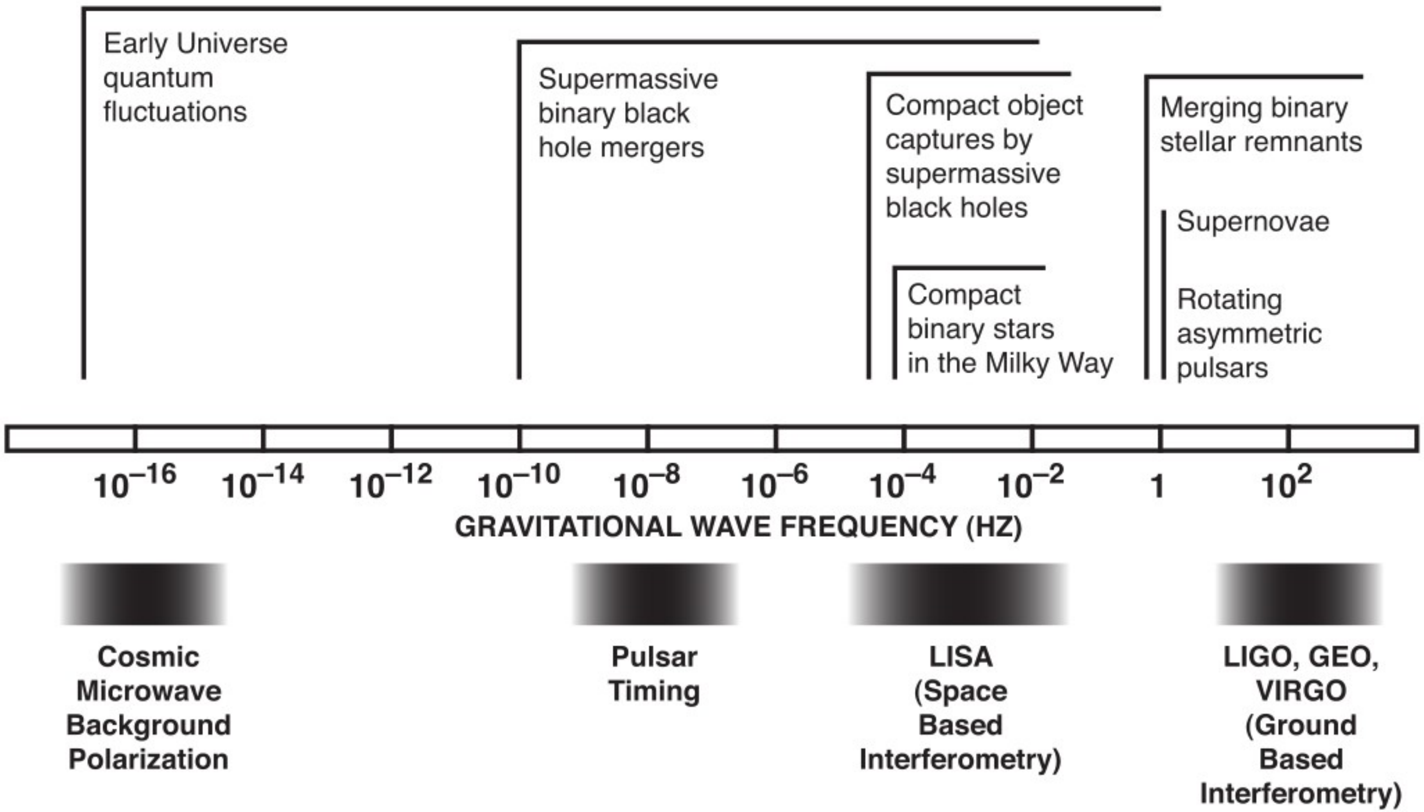


Fig. 3.6 The geometry of the problem in a frame (x', y', z') where a fixed observer is at large distance along the positive z' axis. The normal to the orbit makes an angle ι with the z' axis.

Some Constants

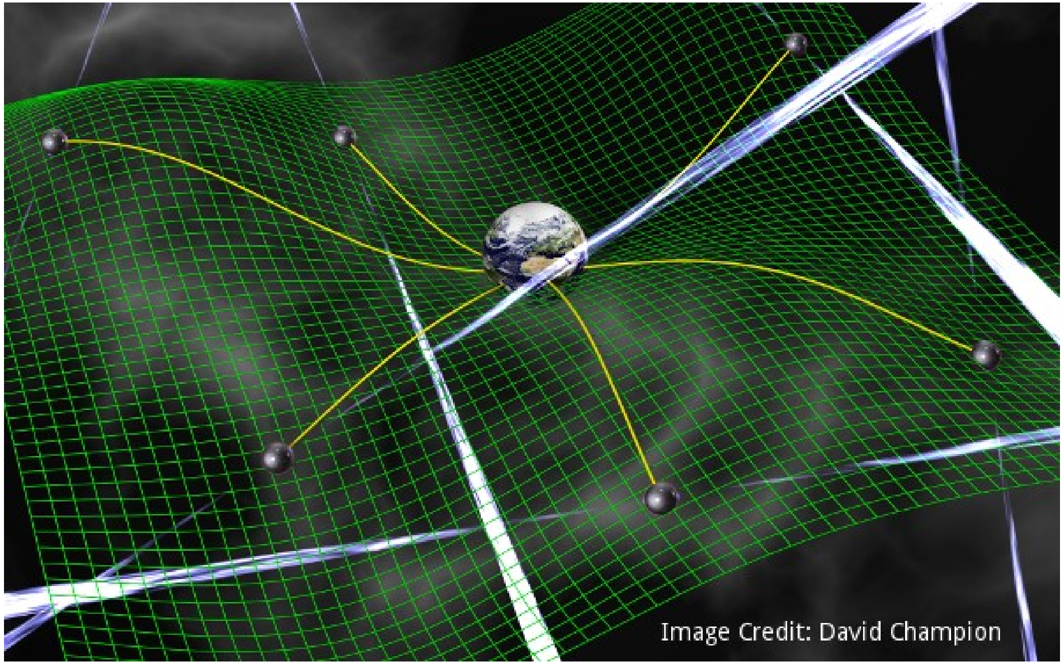
Schwarzschild radius	2GM/c^2	2953 m for Msun
GR Units, Mass	GM/c^2	1477 m for Msun
GR Units, Power	c^5/G	3.628e52 W
GR Units, Energy	c^4/G	1.210e44 J/m
parsec, pc		3.09e16 m = 3.262 ly
astronomical unit, au		149.6e9 m
light-year, ly		0.946e16 m = 0.307 pc
fine structure constant	e^2/(4pi eps0)/hbar c	0.00730 = 1/137.04
wavelength-energy	h c	1.24 eV µm, 1.24 GeV fm

Gravitational Wave Spectrum



Pulsar Timing Arrays

- Millisecond pulsars are very good clocks scattered around the Universe.
- By measuring them once in a while, you can detect a hiccup in the reception of the pulse at Earth. With enough "clocks" you can decide if there was a gravitational wave and from where it may have come.
- See chart, sensitive to very, very long wavelengths, and thus very, very small frequencies. Size of the Universe sort of waves!



Ref: NANOGrav collaboration

Pulsar Timing Arrays

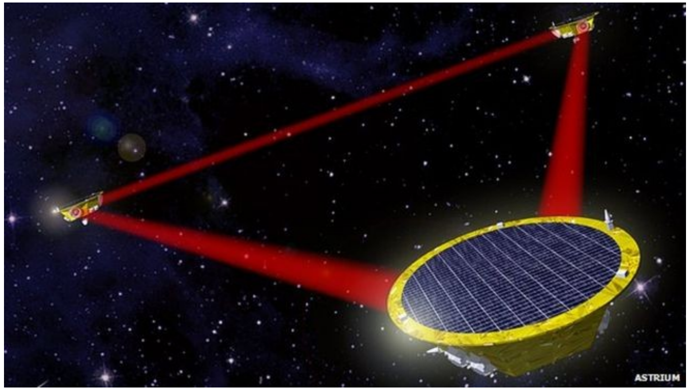
- Several collaborations working at radio astronomy facilities to measure millisecond pulsars:
- NANOGrav, US and NASA, using ~~Arecibo and~~ Green Bank.
- The International Pulsar Timing Array uses many radio telescopes.
- The European Pulsar Timing Array
- The Parkes Pulsar Timing Array (part of the IPTA)
- and others.

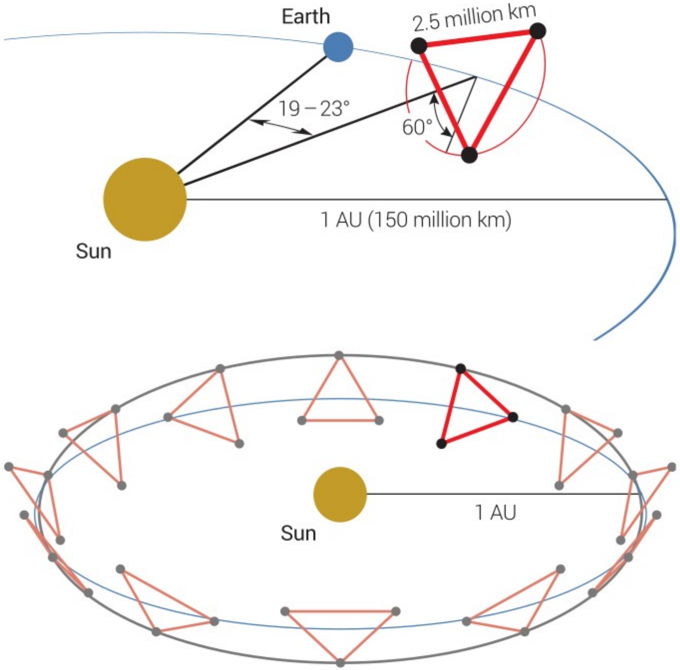
Space-Based

- **Laser Interferometer Space Antenna, LISA**, is a set of satellites with lasers bouncing in a triangle arrangement.
  - No seismic effects!
  - Very long baseline, 2.5e9 m each leg.
  - Control of craft subtle, “drag-free flying” with free test masses inside as reference.
- Currently mostly European LISA Mission with NASA/US support. LISA could launch in 2037 (factsheet).
  - In the hardware testing/development and design phase, soon to choose the aerospace company that will build the spacecraft.
- LISA Pathfinder satellite experiment tested several components for eLISA, early 2016.

LISA

- LISA has a full equilateral triangle of lasers and baseline of 2.5e9 m. ESA mission with NASA participation.

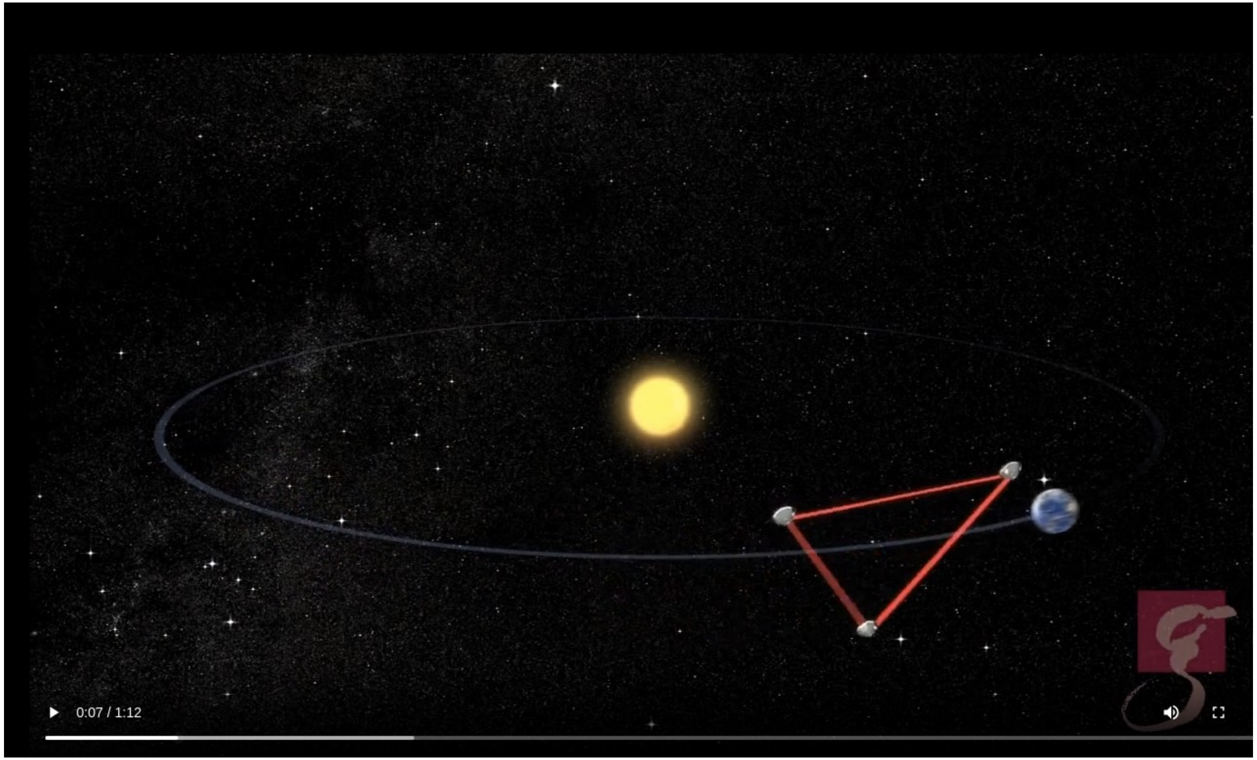




- Chinese have two groups working on this idea:
  - TianQin
  - TAIJI

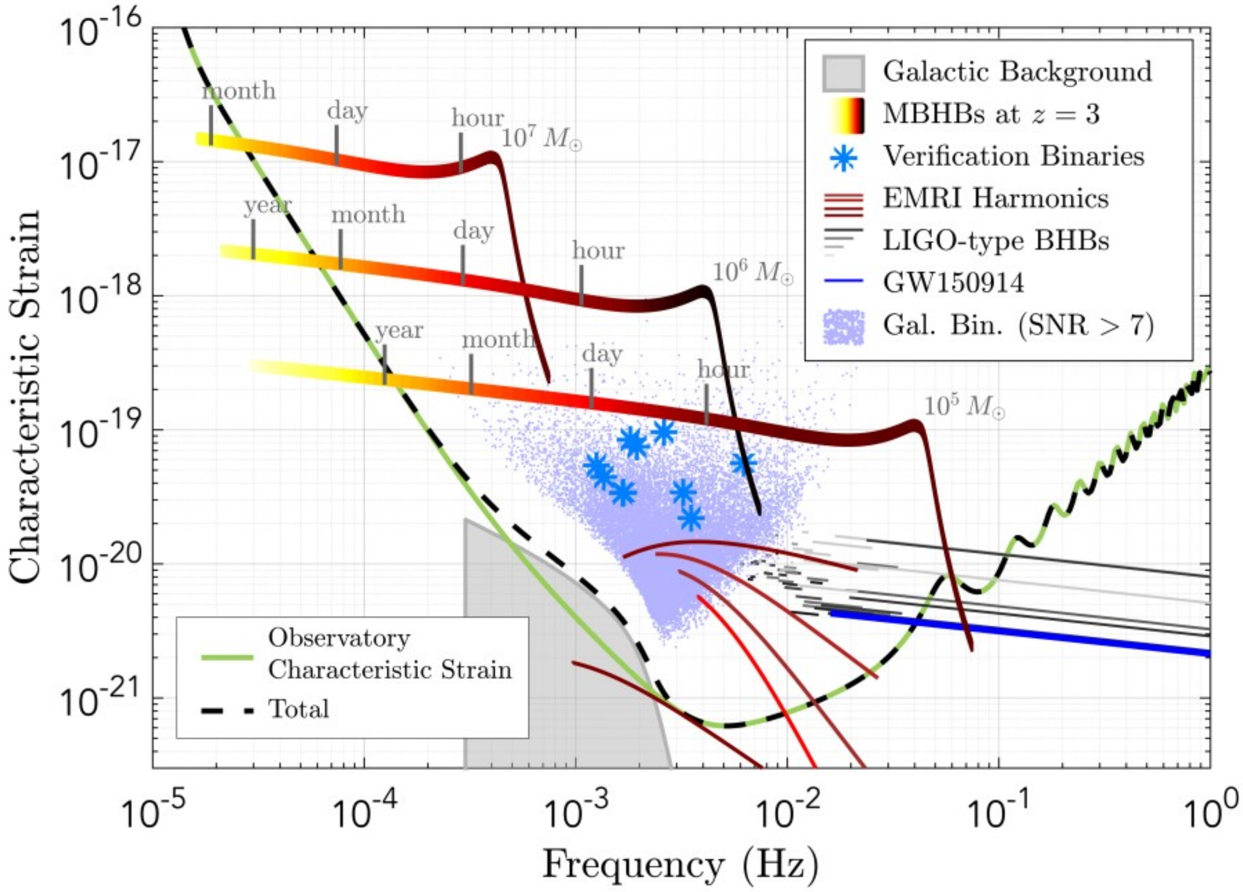
Credit

LISA Orbit



https://svs.gsfc.nasa.gov/vis/a030000/a030500/a030569/slide_09-A1_LISA_orbits2.mp4

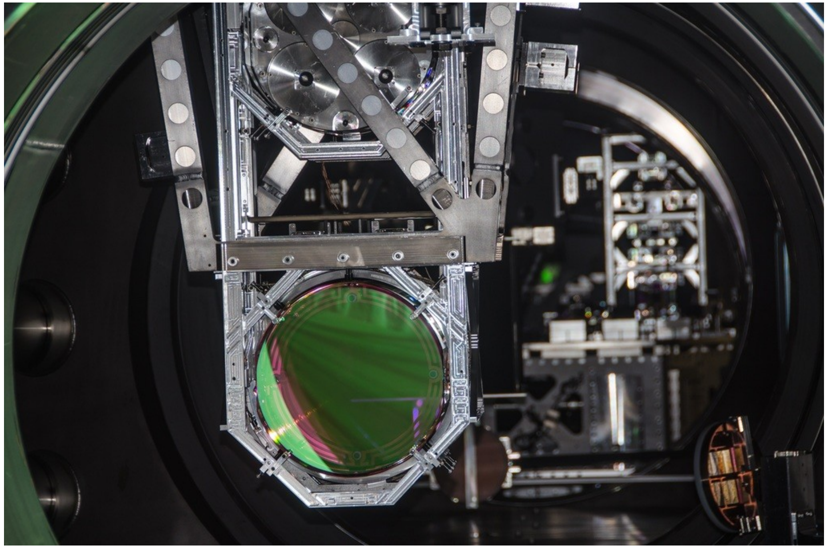
LISA Astrophysical Sources

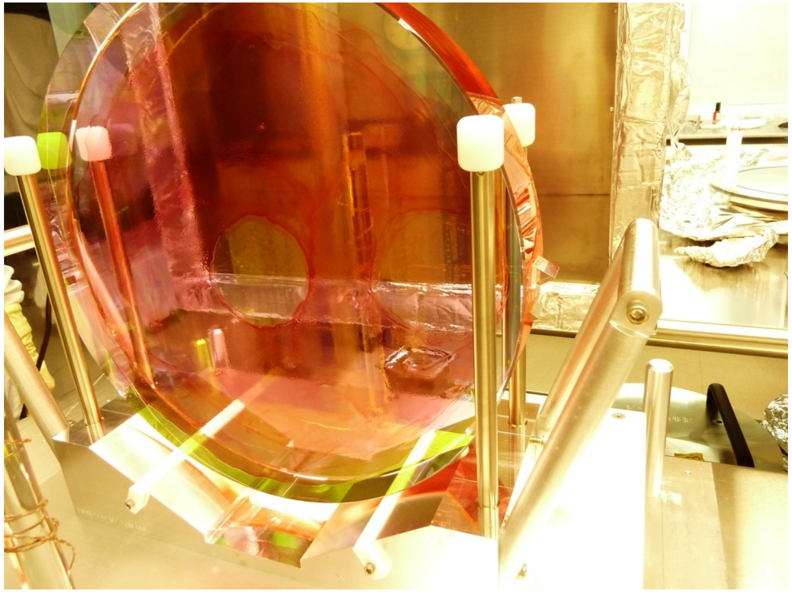


Credit

LIGO Hardware

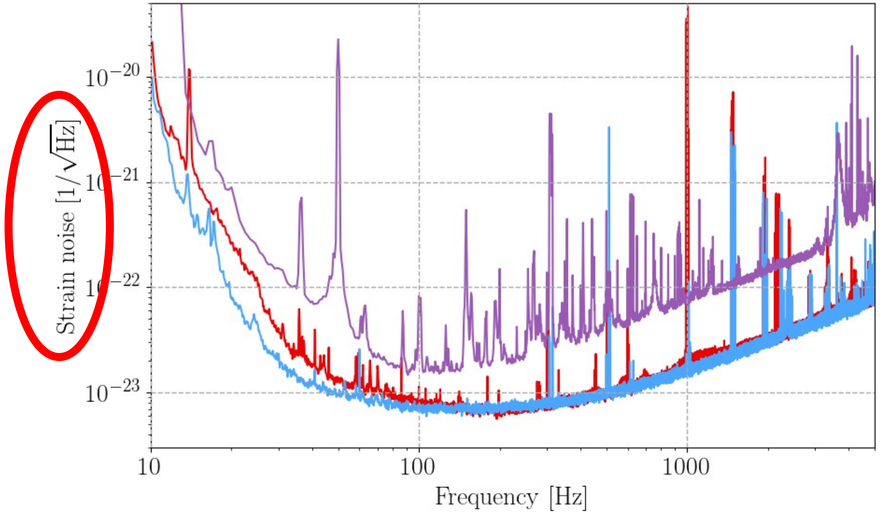
- LIGO “Test Mass” in suspension

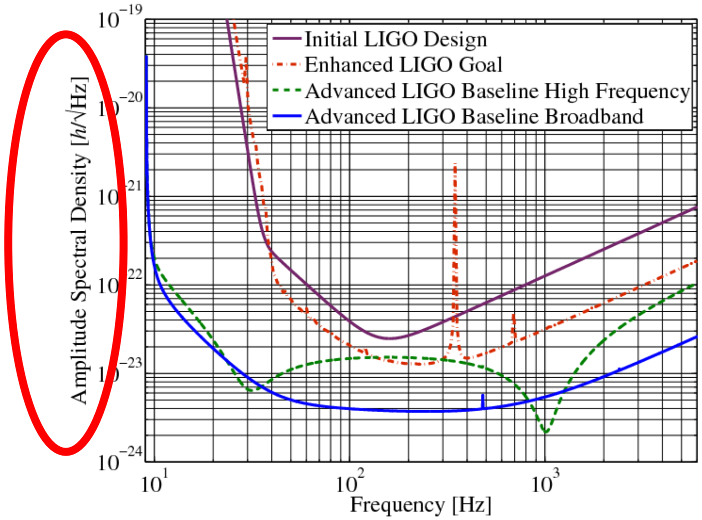




- Beam Splitter

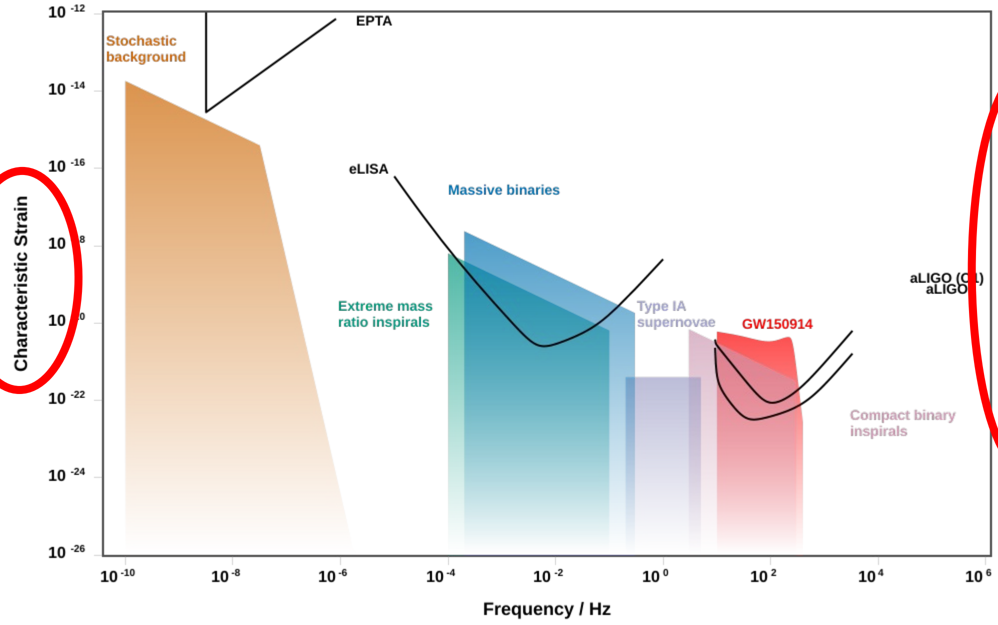
Strain Curve, Strain Sensitivity, Amplitude Spectral Density, Power Spectral Density...OH MY!

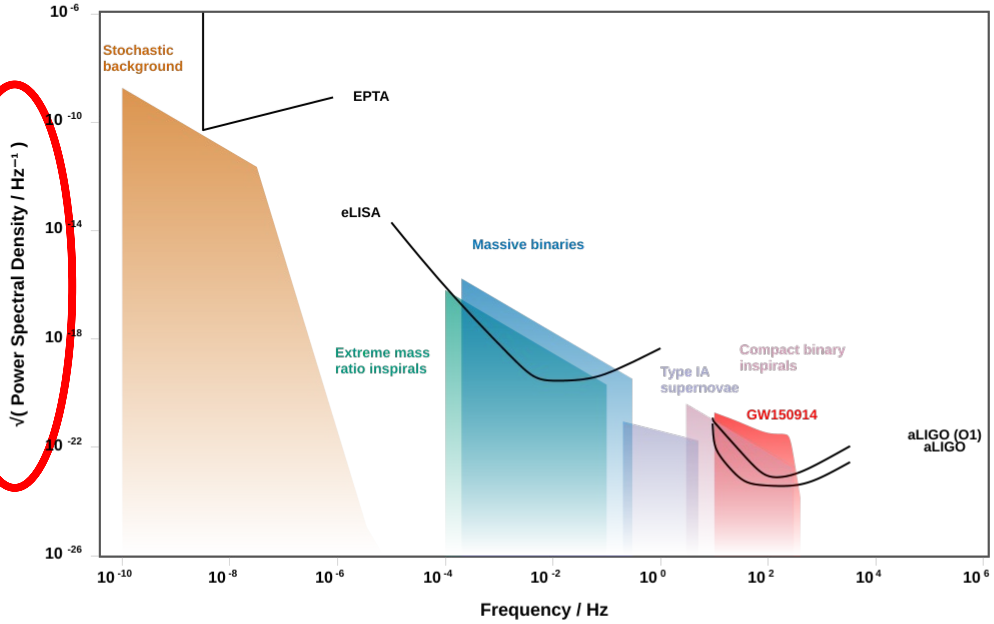




“per root Hz” is usually called the *strain sensitivity*, or sometimes *amplitude spectral density*

Strain Curve, etc…





Characteristic Strain

Square root of *Power Spectral Density* is the *amplitude spectral density.*

Ref: Cole's online app .