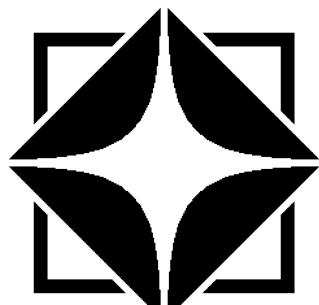


REAL-TIME RFI MITIGATION FOR THE BEAMFORMER MODE OF THE UPGRADED GIANT METREWAVE RADIO TELESCOPE (GMRT)

**Aditya Chowdhury
Yashwant Gupta**



NCRA • TIFR

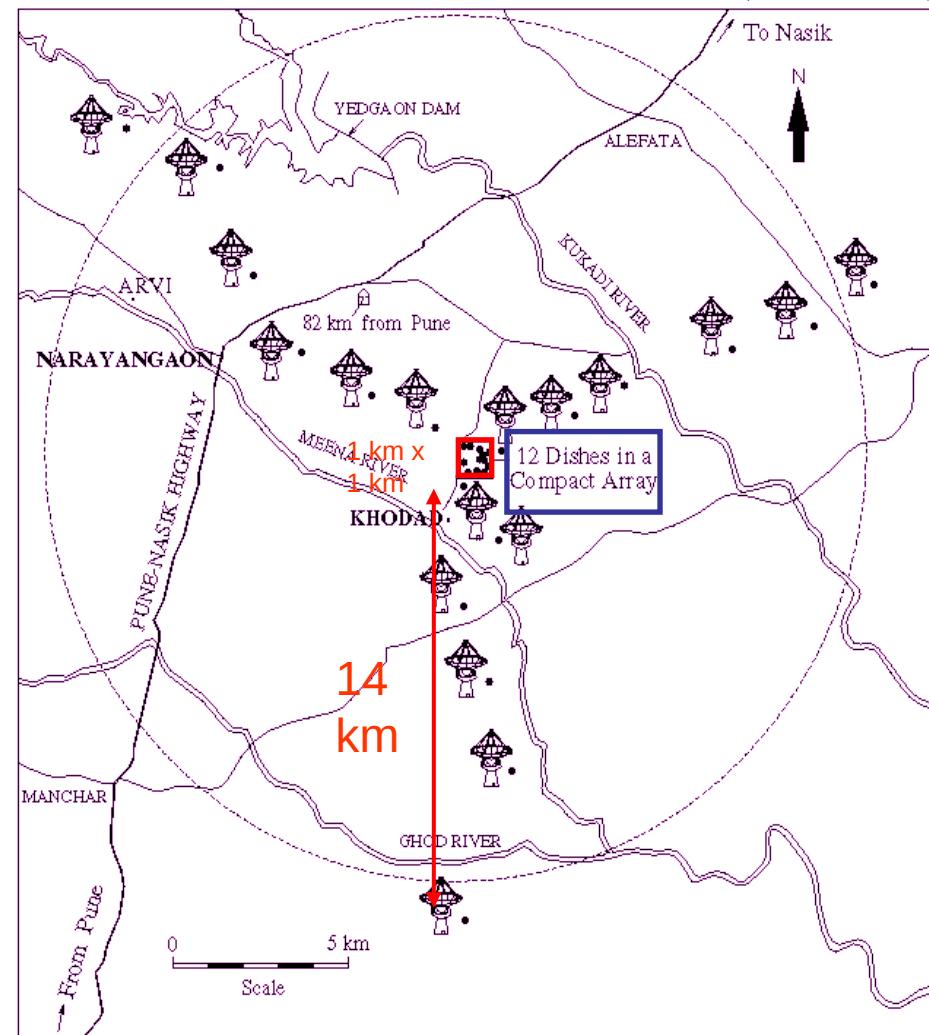
**NATIONAL CENTRE FOR RADIO ASTROPHYSICS
TATA INSTITUTE OF FUNDAMENTAL RESEARCH
Pune, India.**

REAL-TIME RFI MITIGATION FOR THE BEAMFORMER MODE OF THE UPGRADED GIANT METREWAVE RADIO TELESCOPE (GMRT)



- 30 dishes, 45 m diameter each
 - 12 dishes in a central 1 km x 1 km region (central square)
 - remaining along 3 arms of Y-shaped array
 - baselines : ~ 200 m (shortest);
~ 30 km (longest)
- Frequency range :
 - **130-170 MHz**
 - **225-245 MHz**
 - **300-360 MHz**
 - **580-660 MHz**
 - **1000-1450 MHz**
 - **max instantaneous BW = 32 MHz**
- Effective collecting area (2-3% of SKA) :
 - 30,000 sq m at lower frequencies
 - 20,000 sq m at highest frequencies
- Supports 2 modes of operation :
 - Interferometry, aperture synthesis
 - Array mode (incoherent & coherent)

LOCATIONS OF GMRT ANTENNAS (30 dishes)

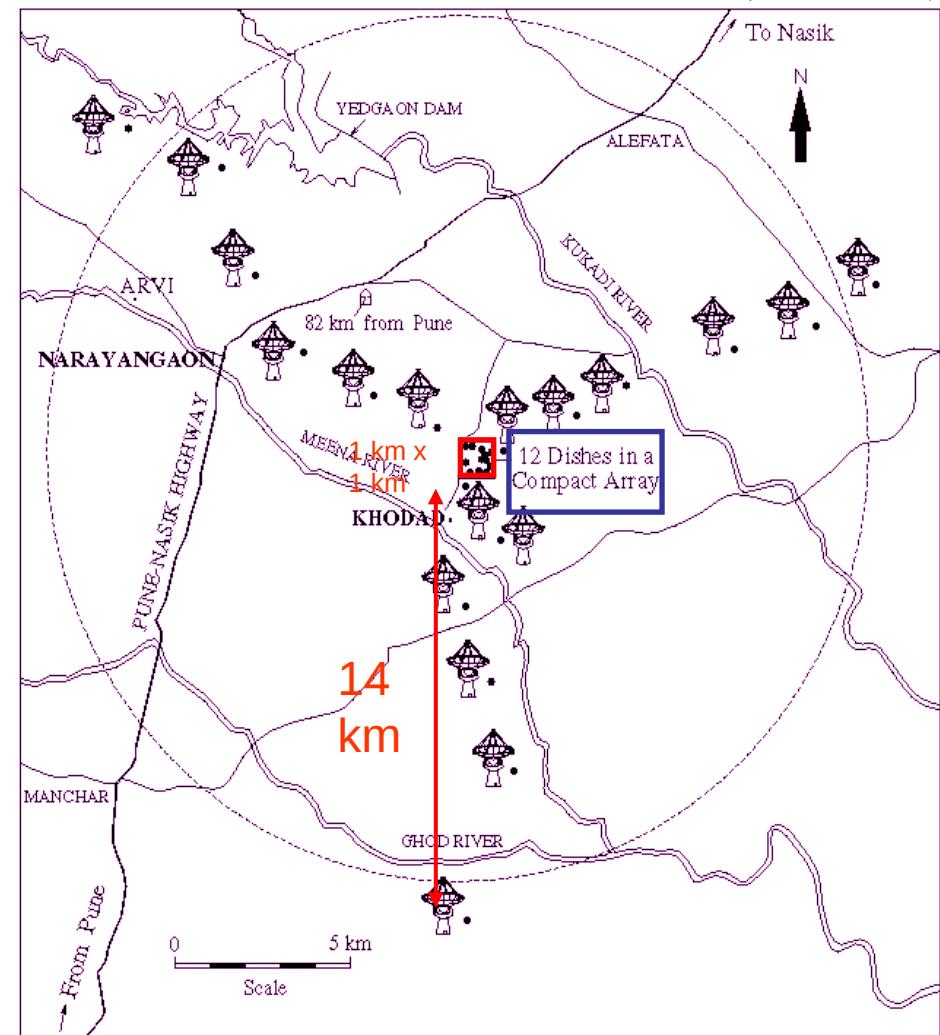


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 - baselines : ~ 200 m (shortest);
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- Frequency range :
 - **125-250 MHz**
 - **250-500 MHz**
 - **550-850 MHz**
 - **1050-1450 MHz**
 - **max instantaneous BW = 400 MHz**
- Effective collecting area (2-3% of SKA) :
 - 30,000 sq m at lower frequencies
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- Supports 2 modes of operation :
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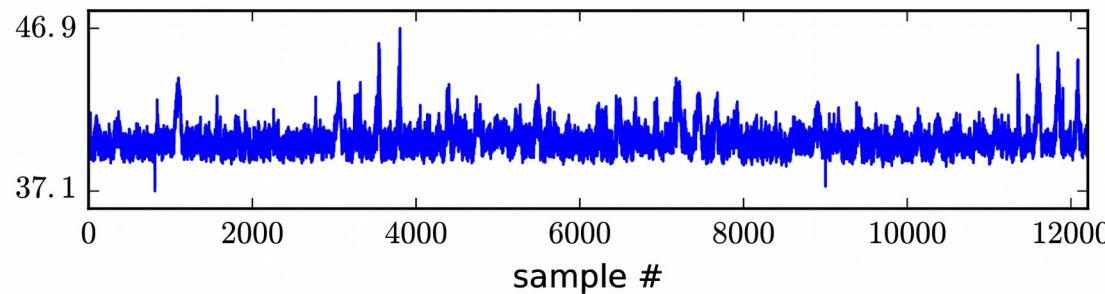
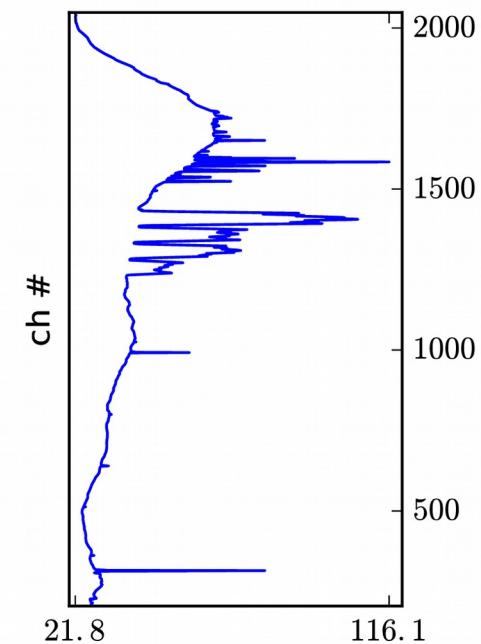
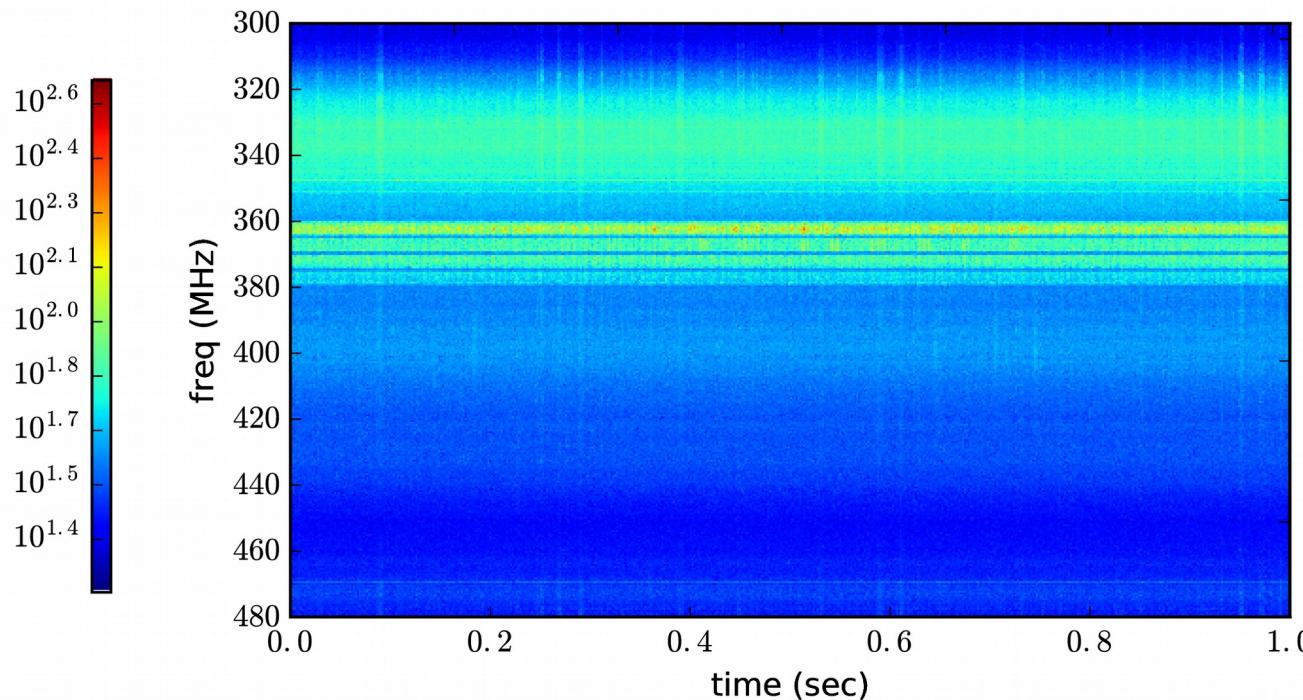
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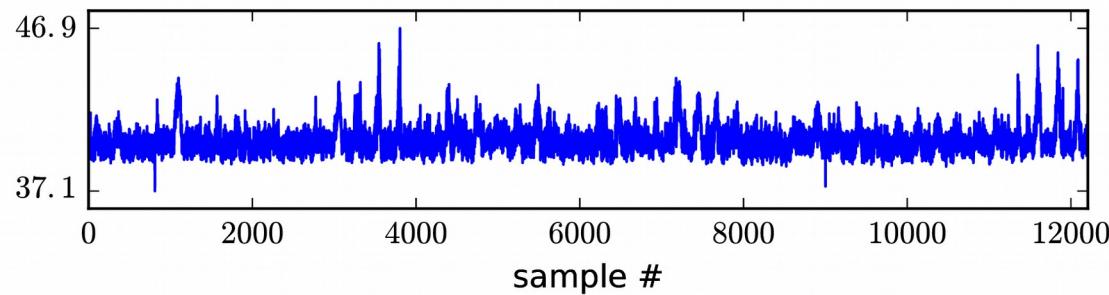
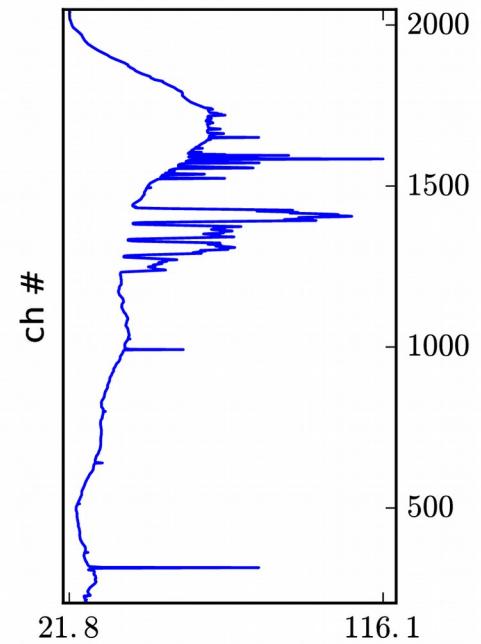
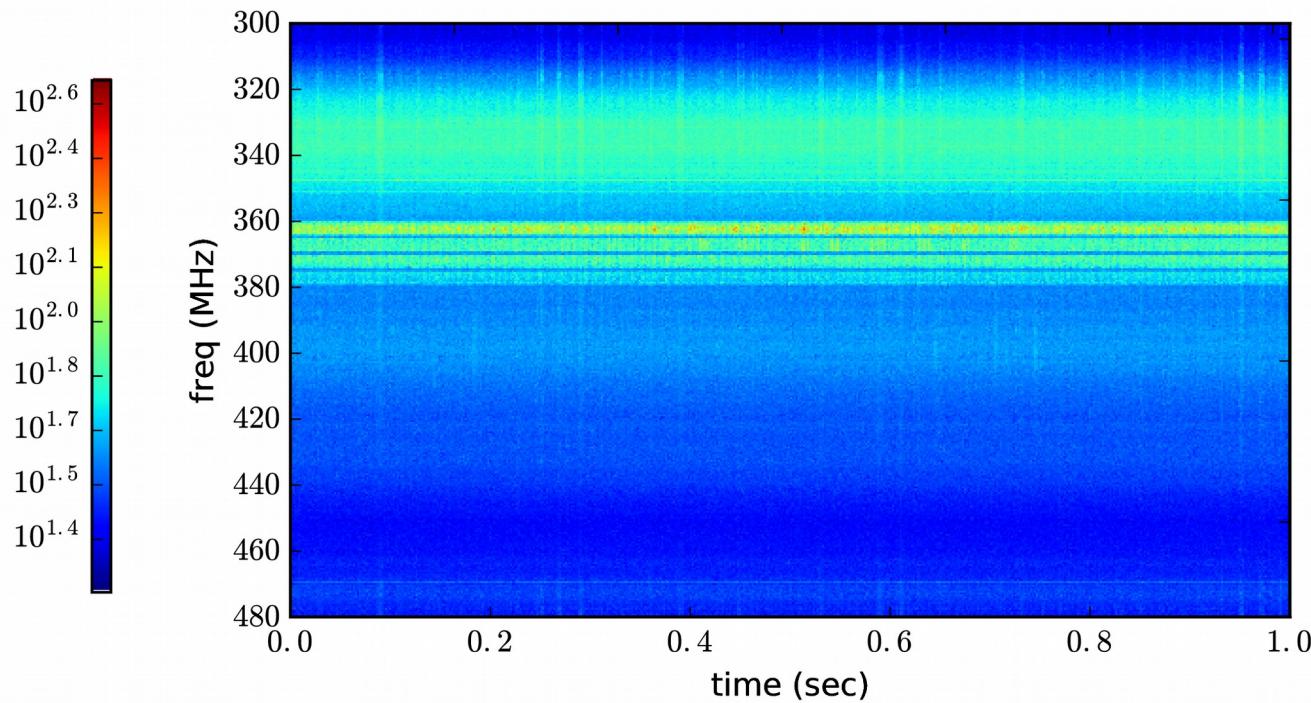
REAL-TIME RFI MITIGATION FOR THE BEAMFORMER MODE of THE UPGRADED GIANT METREWAVE RADIO TELESCOPE (GMRT)



- Signals from all antennas added either coherently or incoherently.
- Used mostly to study unresolvable objects - pulsars and transients.



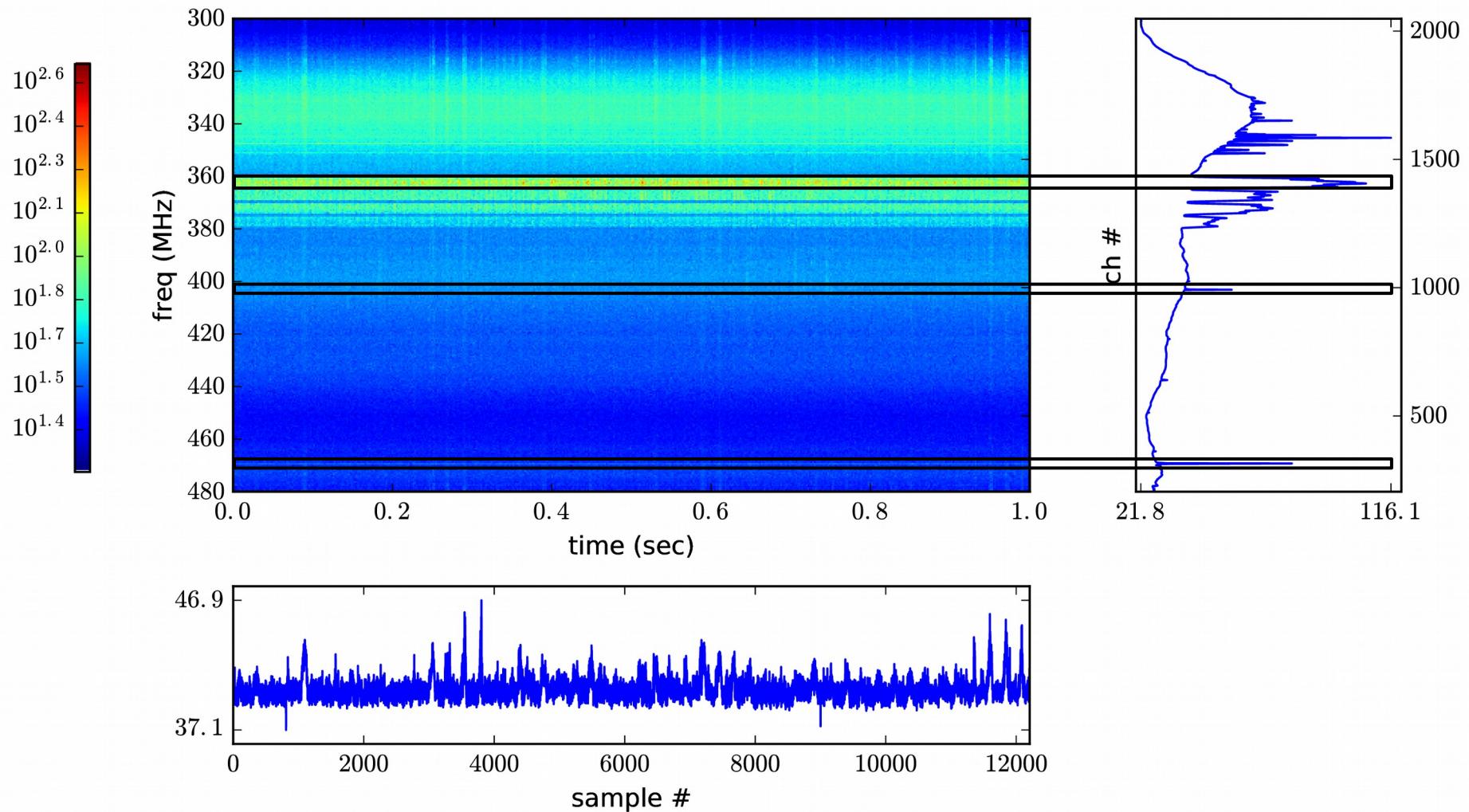
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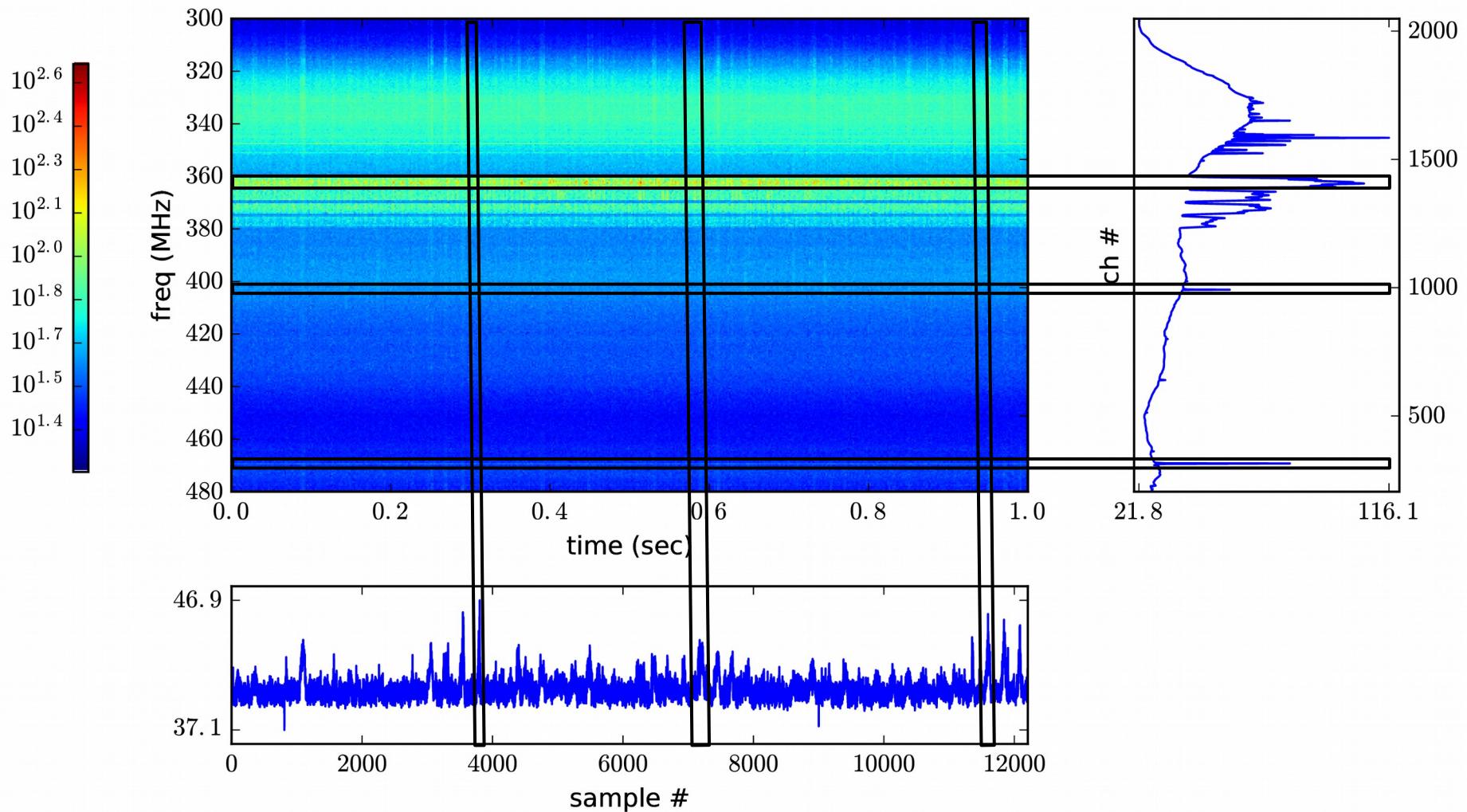
→ “Narrow” band



REAL-TIME RFI MITIGATION FOR THE BEAMFORMER MODE OF THE UPGRADED GIANT METREWAVE RADIO TELESCOPE (GMRT)



- “Narrow” band
- “Wide” band



Detecting Outliers in a Series



Objective: *An algorithm that can detect the “true” mean and std. dev. of N samples derived from a gaussian distribution with $p\%$ corrupted data. Usual mean and RMS is a very rapid function of p .*

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- Median is taken as the estimate of mean.
- Median of Absolute Deviation from the median is used to estimate std. dev.

$$\sigma = 1.4826 \text{median}_i(|X_i - \text{median}_j(X_j)|)$$

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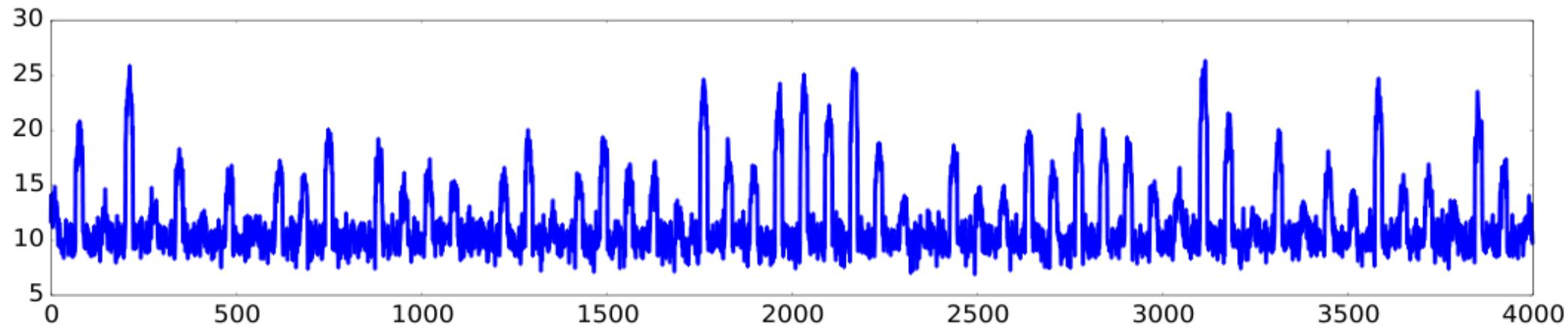
- Computationally O(n log n) because of a sort operation.

- **Can we do better? Certainly!**

Detecting Outliers in a Series : Smarter than MAD



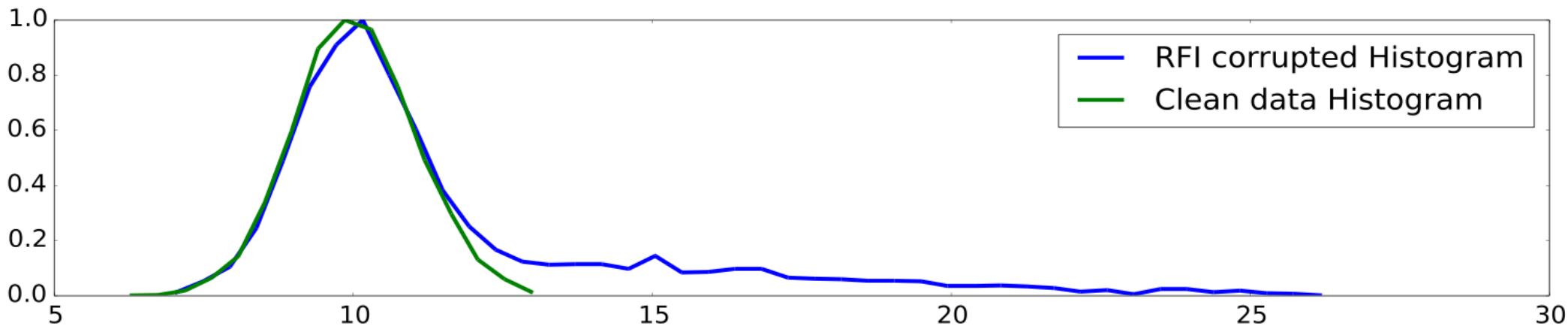
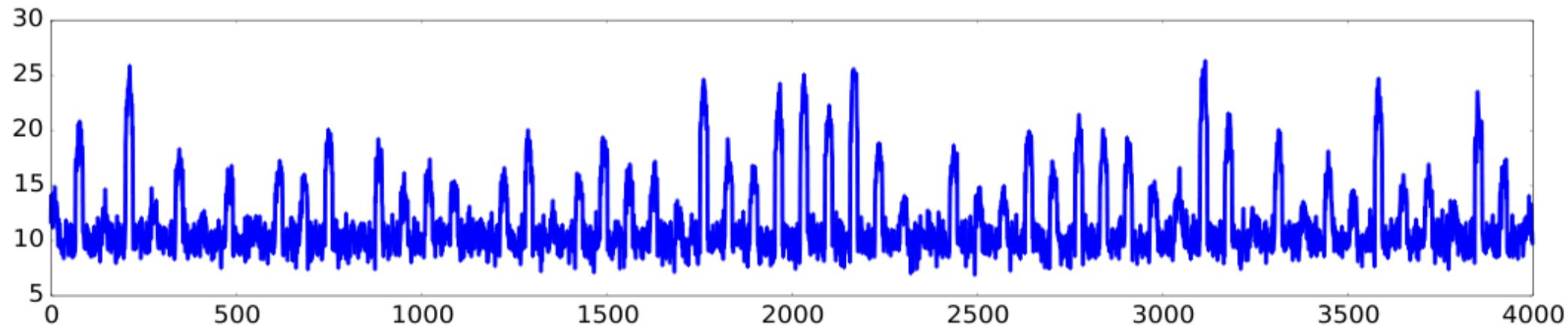
A simulated time series with 30% corruption:



Detecting Outliers in a Series : Smarter than MAD



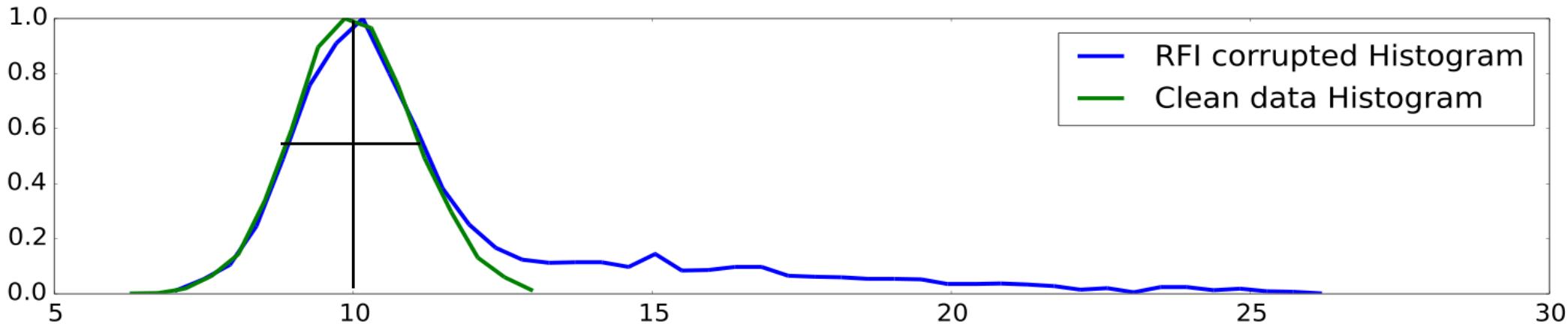
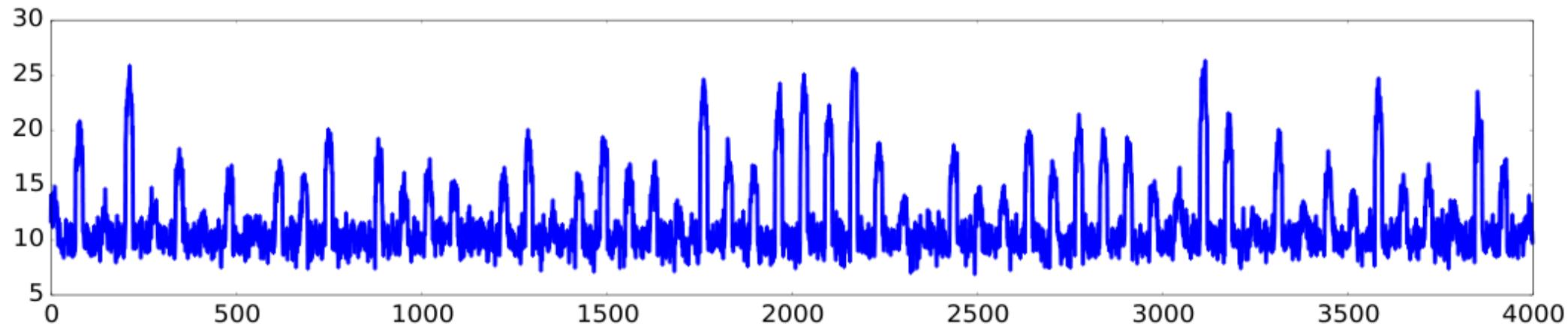
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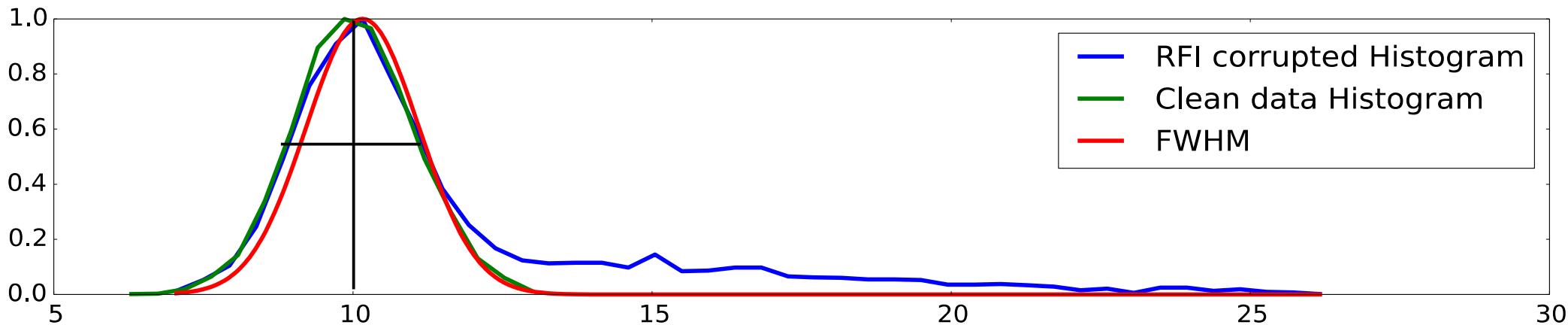
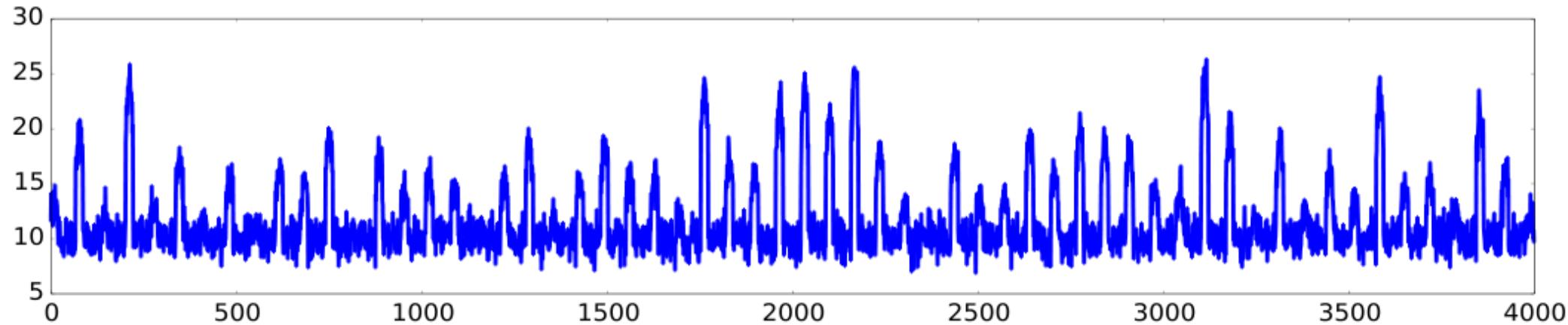
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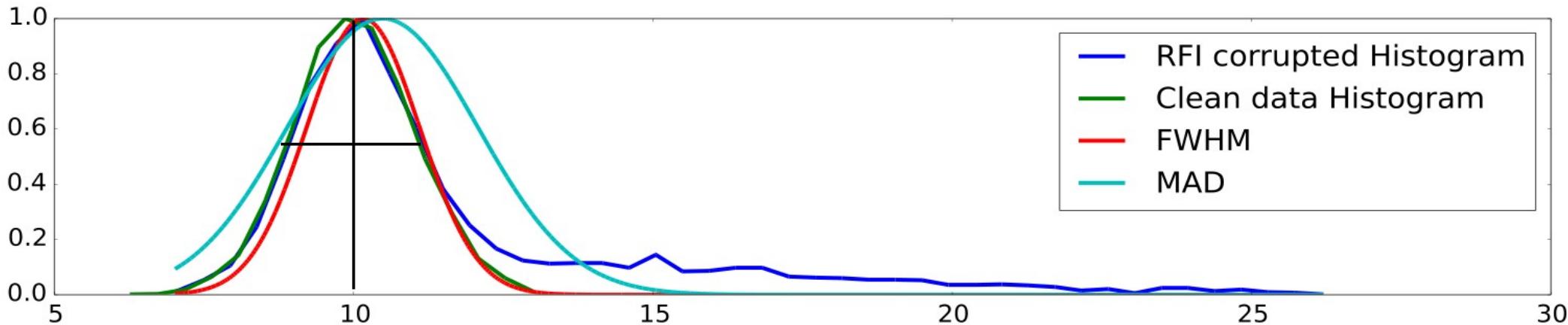
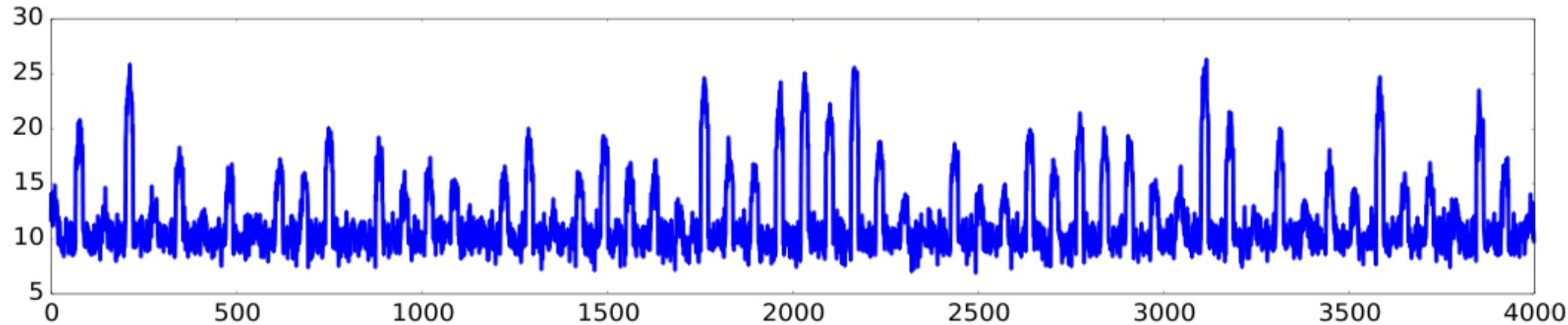
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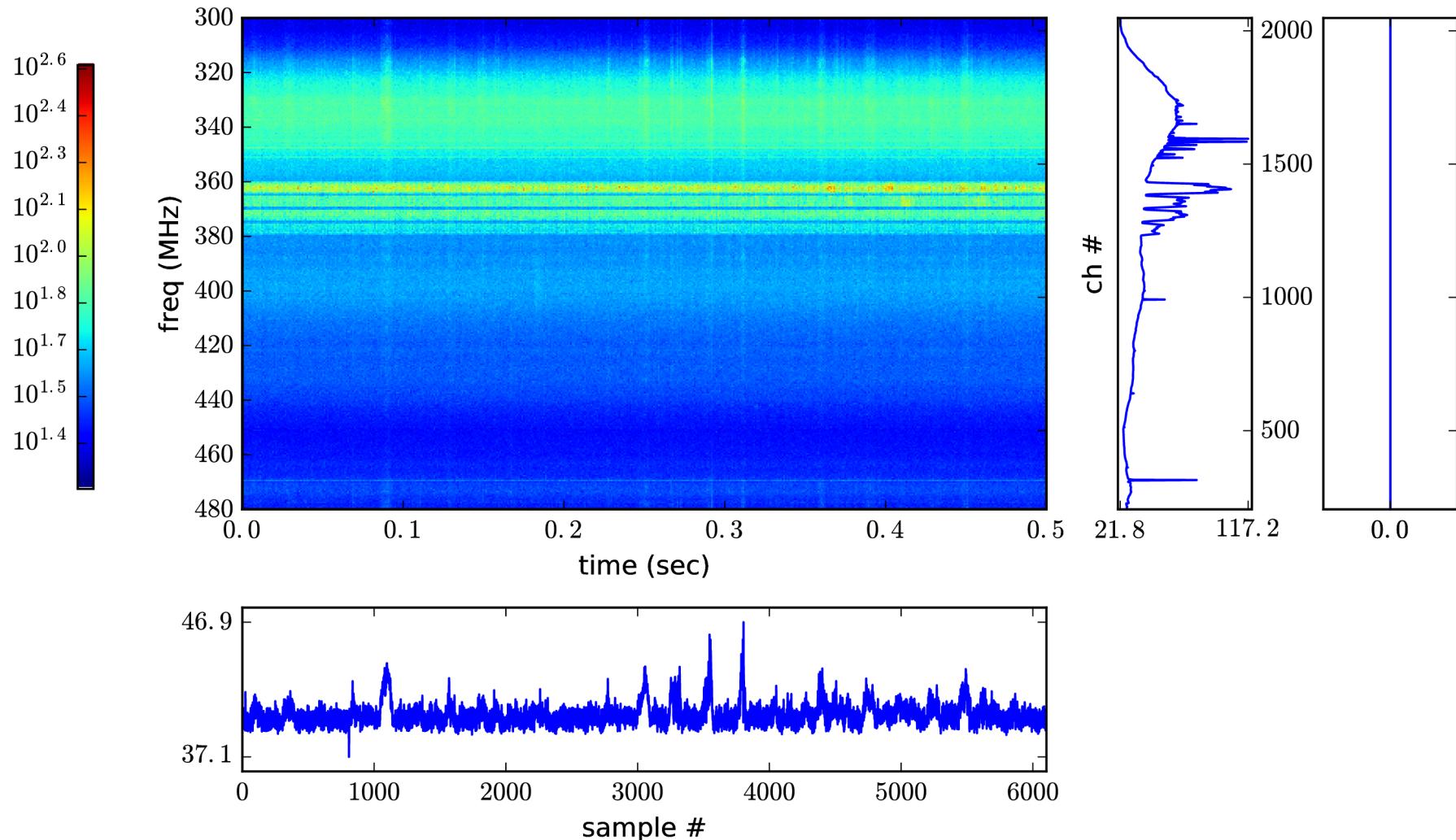
A simulated time series with 30% corruption:



Detecting RFI in Beamformer Data



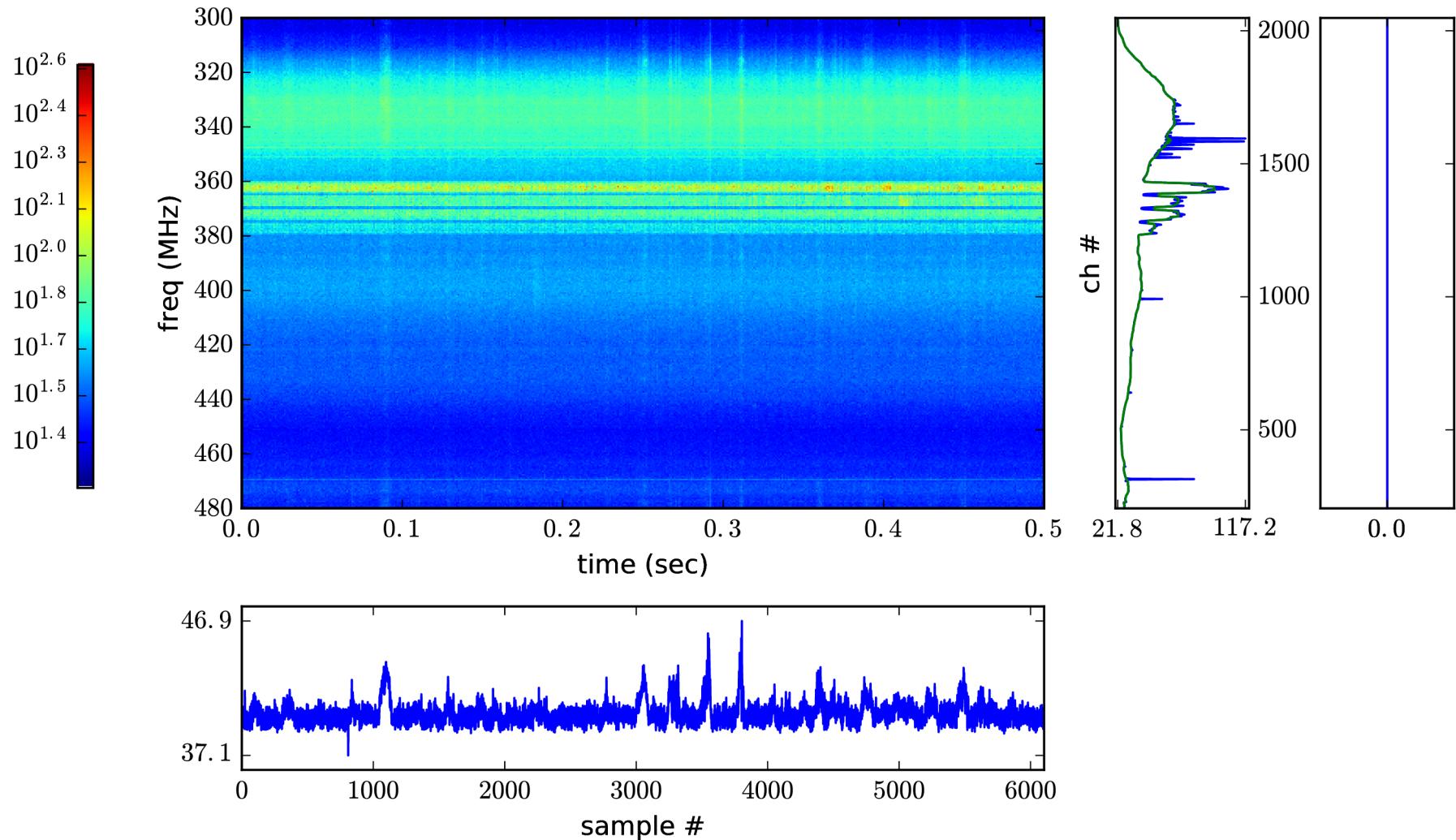
→ Beamformer Data



Detecting RFI in Beamformer Data



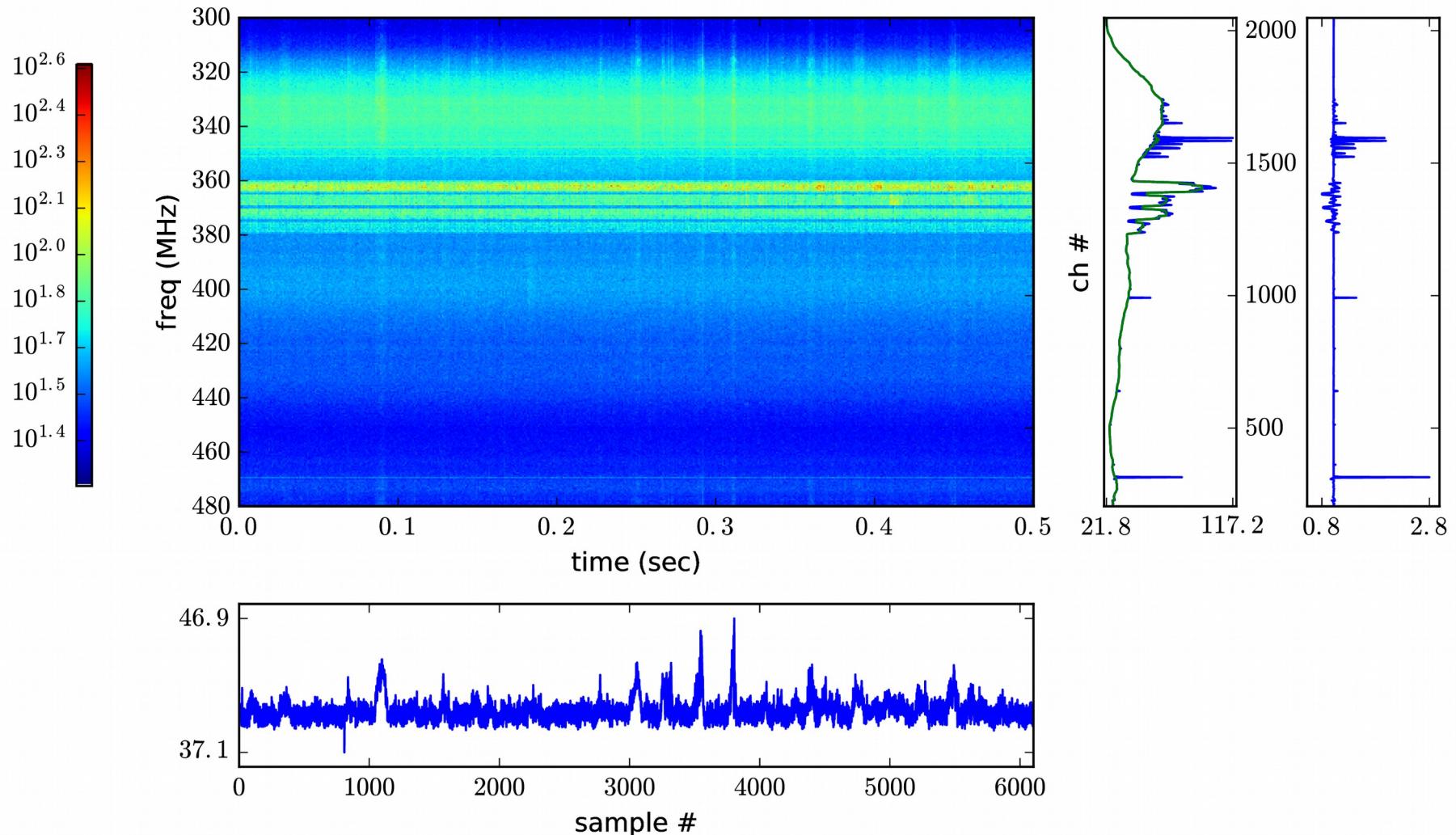
- Detecting Narrow-band RFI: Smooth bandshape.



Detecting RFI in Beamformer Data



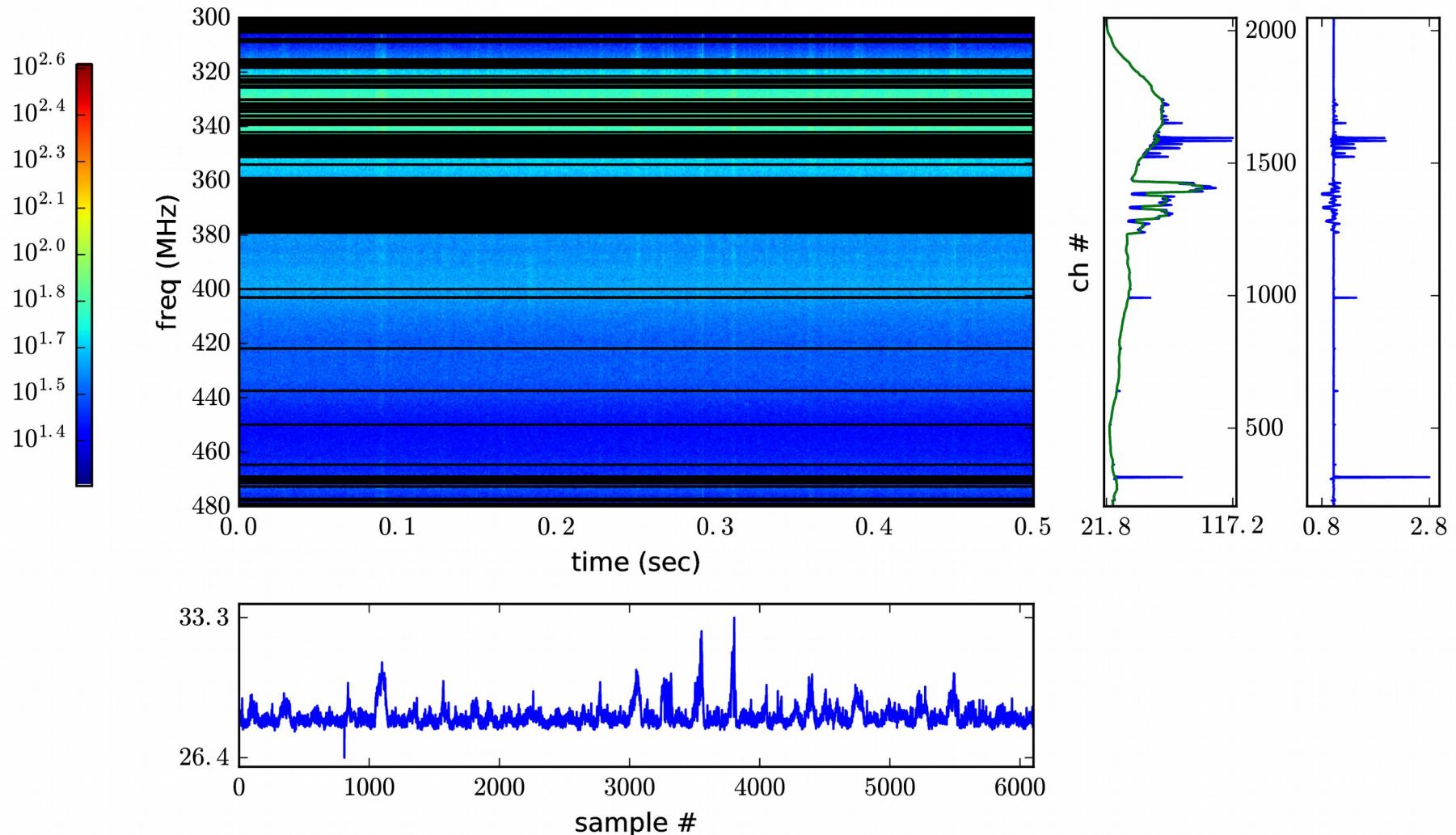
- Detecting Narrow-band RFI: Get power normalized bandshape.



Detecting RFI in Beamformer Data



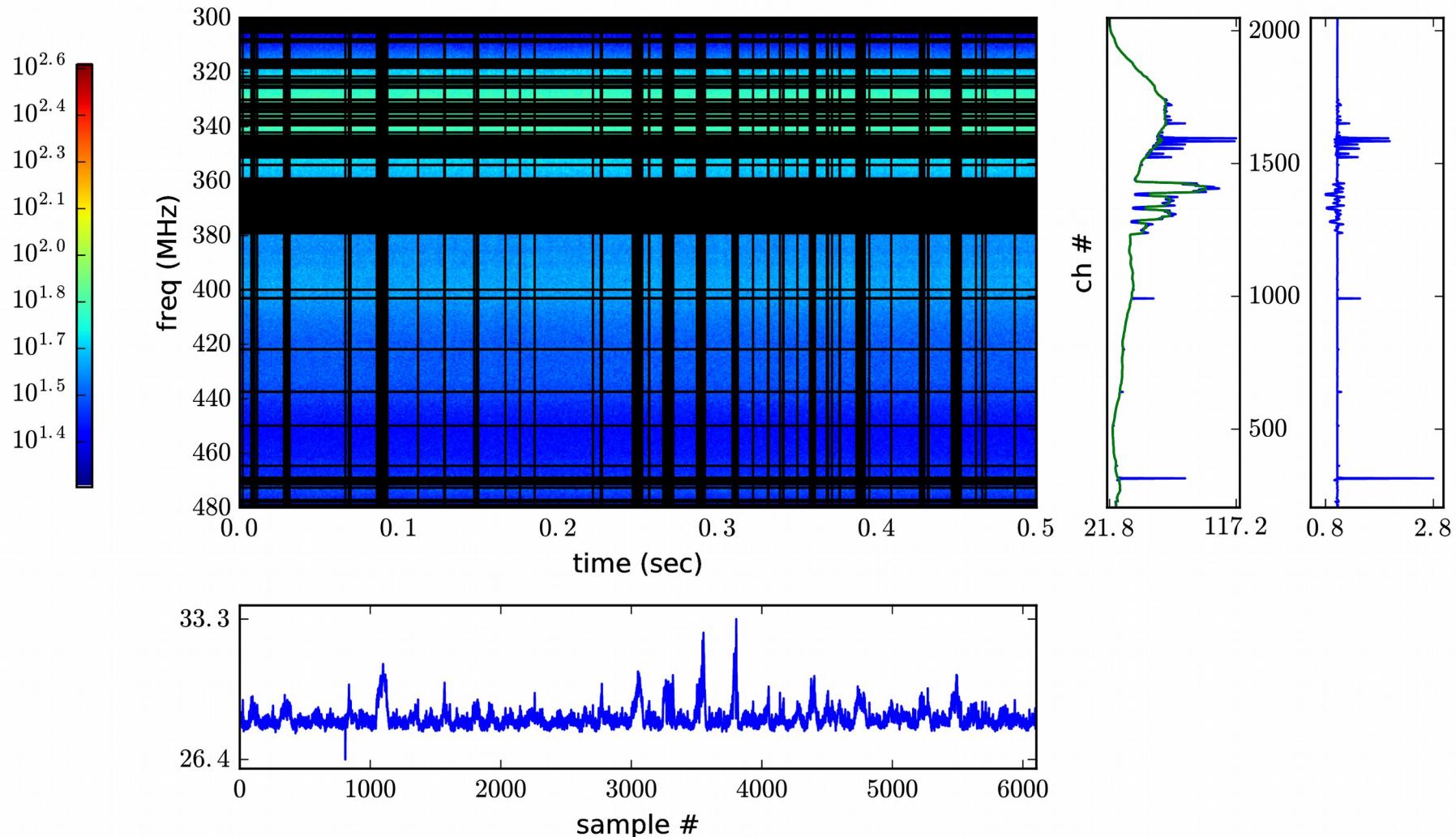
- Detecting Narrow-band RFI: Outlier detection on power normalized bandshape.



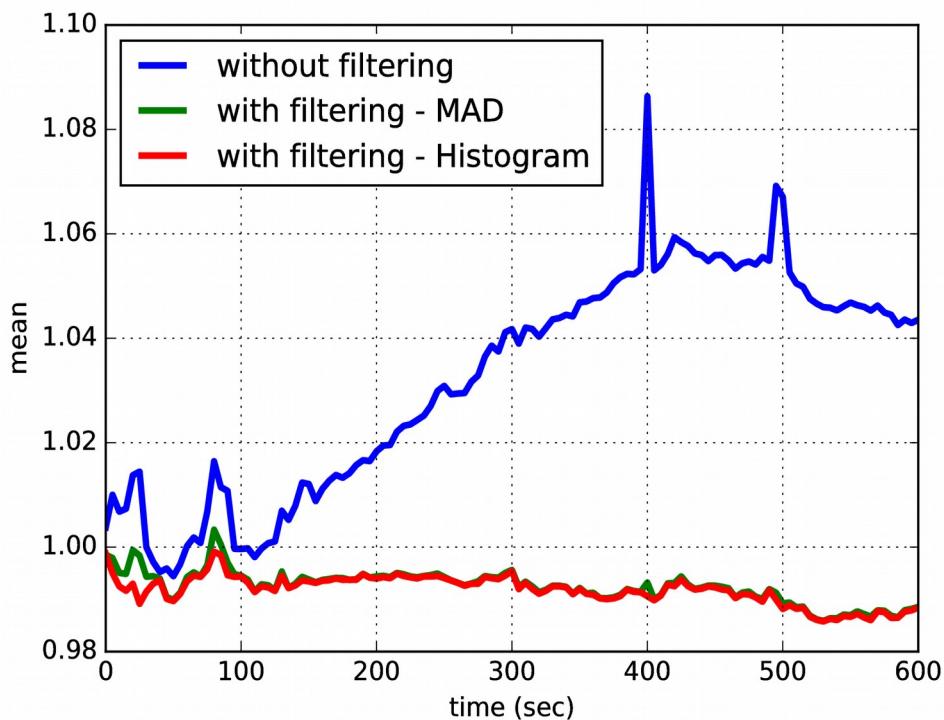
Detecting RFI in Beamformer Data



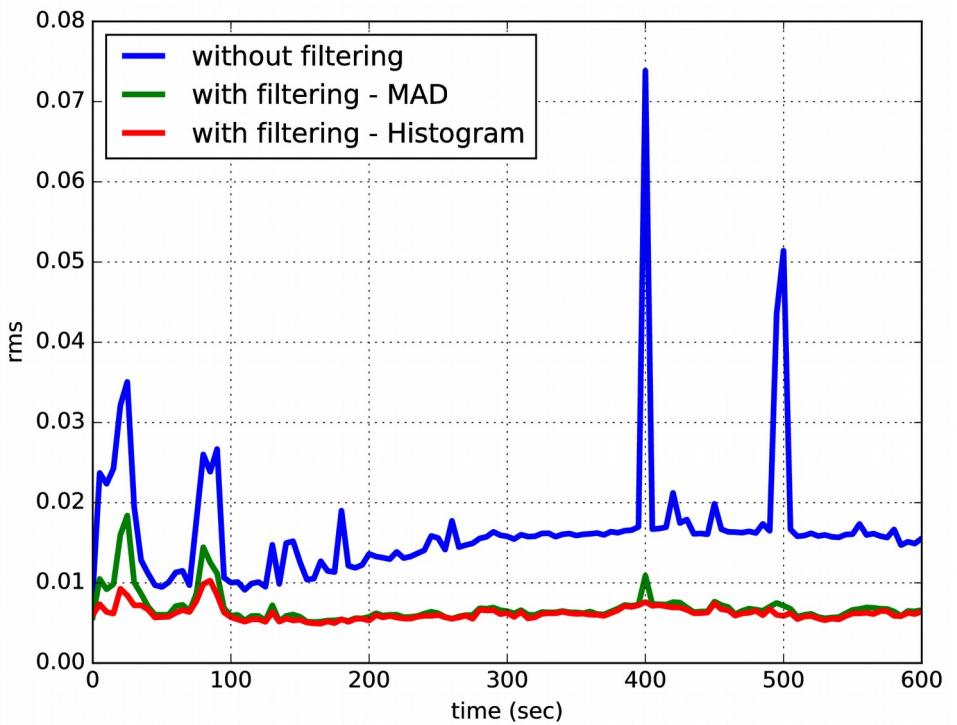
- Detect broad-band RFI from resultant time-series



Results - Stability of Mean and RMS



Mean Variation

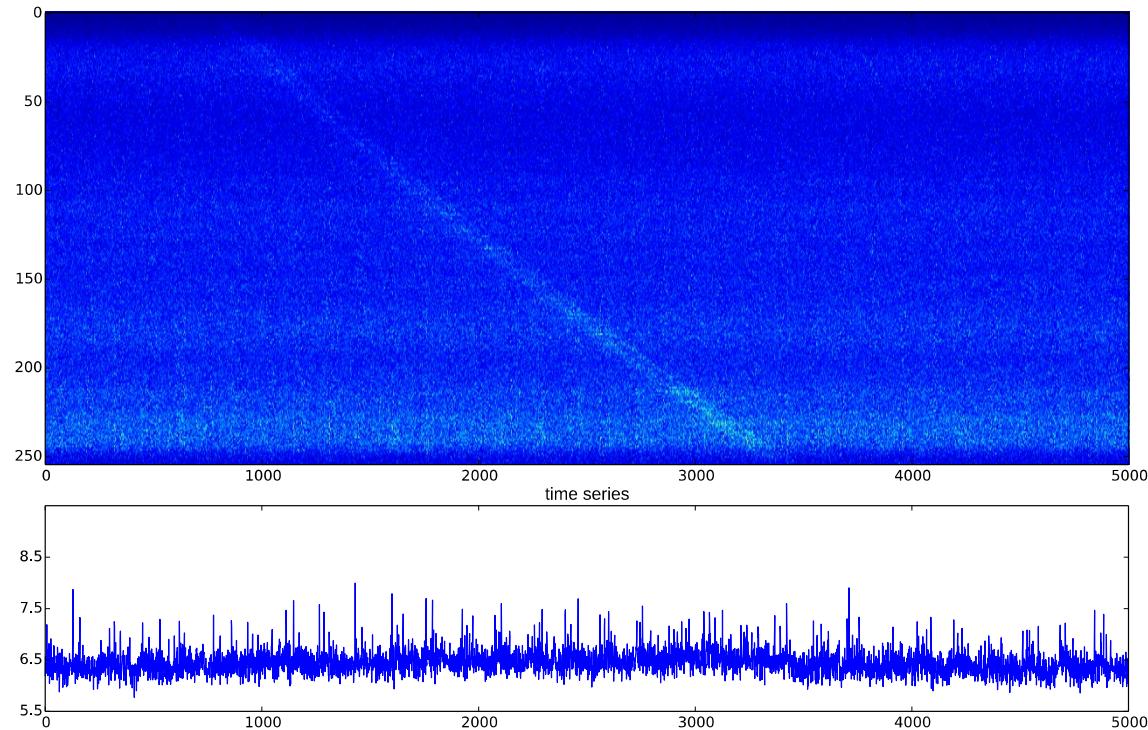


Mean Variation

Introduction to Pulsar signal processing



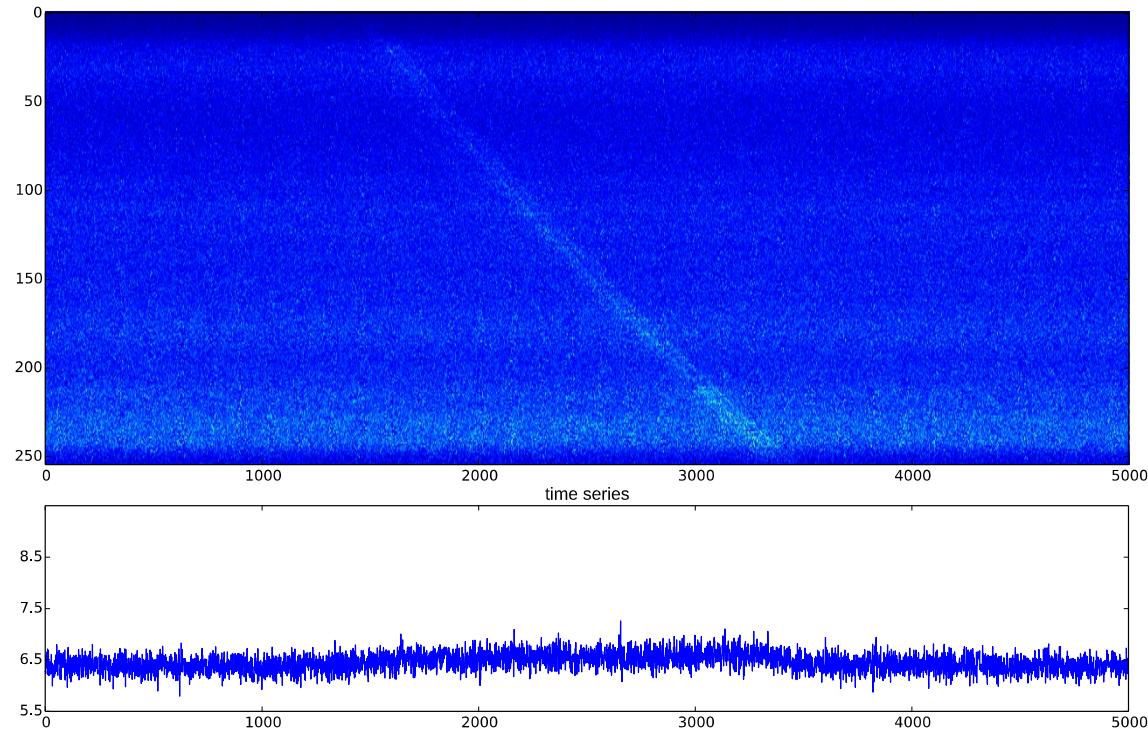
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- Interstellar medium causes dispersion of pulse
- Different time of arrival of the pulse between two frequency channels.
- Needs to be corrected for to prevent smearing.
- For most pulsars, signal buried in noise in uncorrected time series.



Introduction to Pulsar signal processing



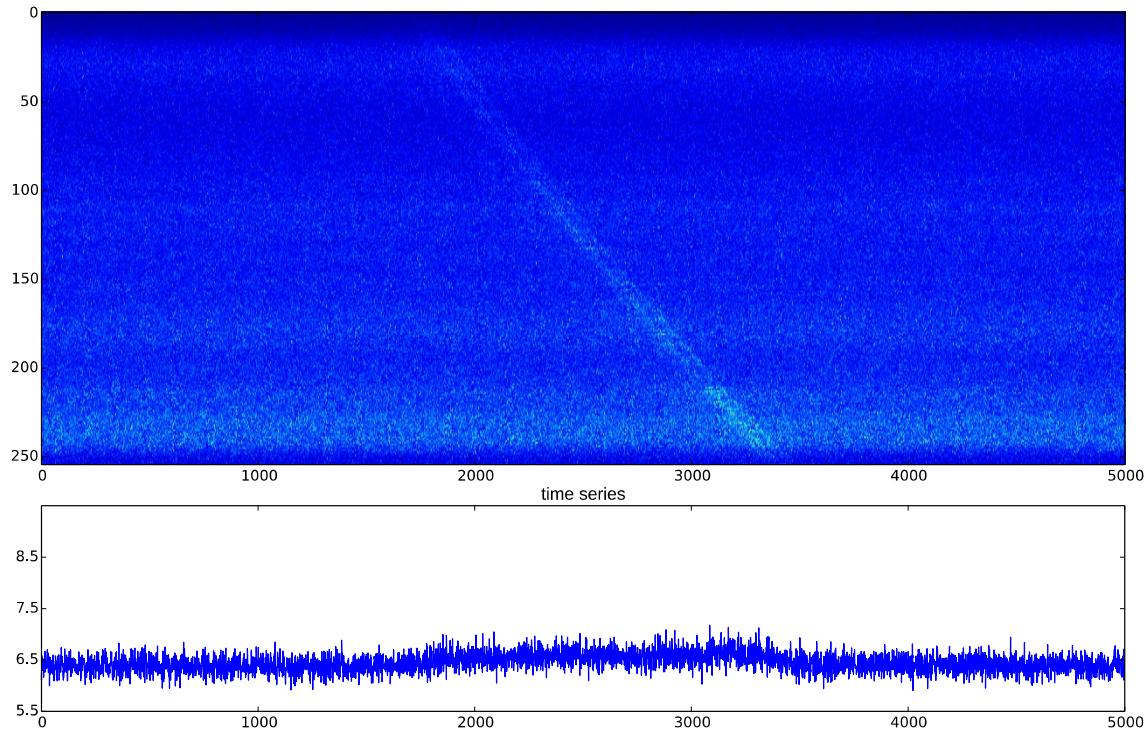
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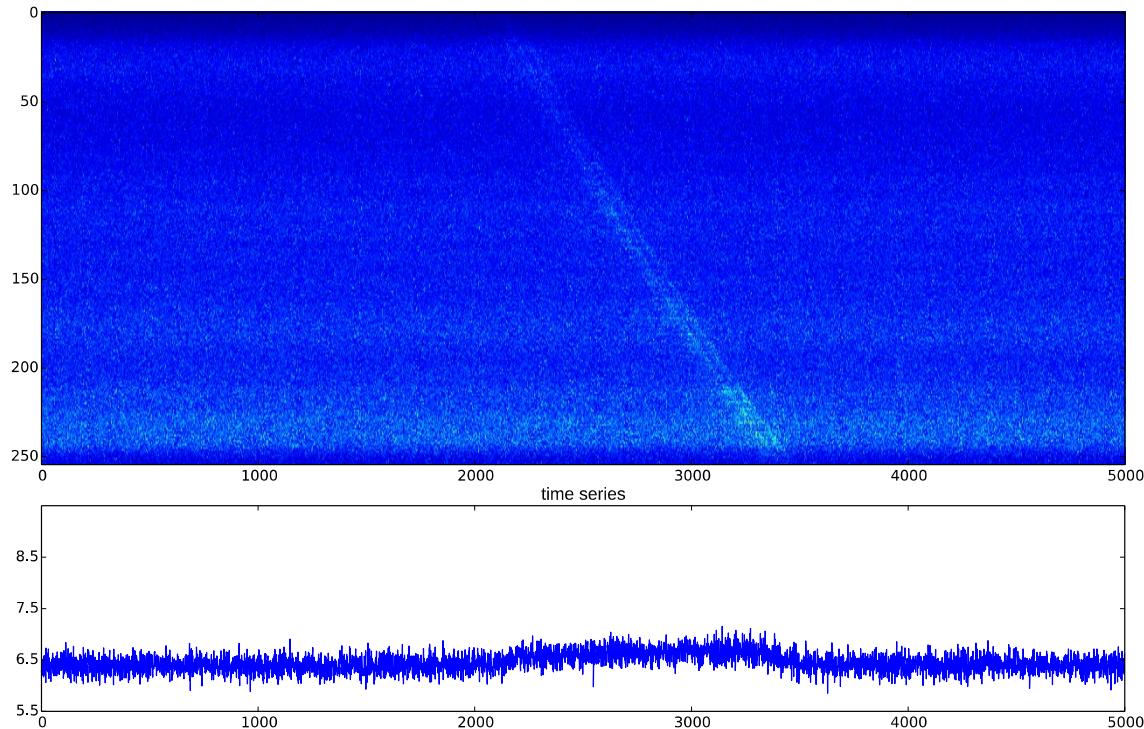
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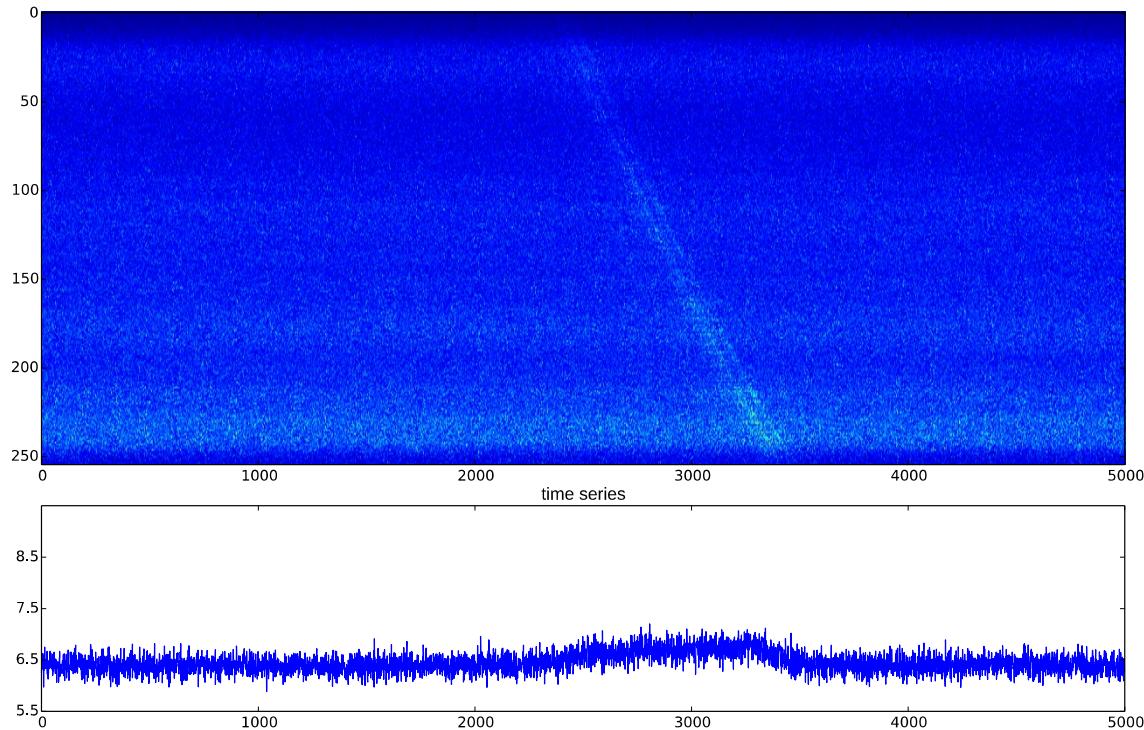
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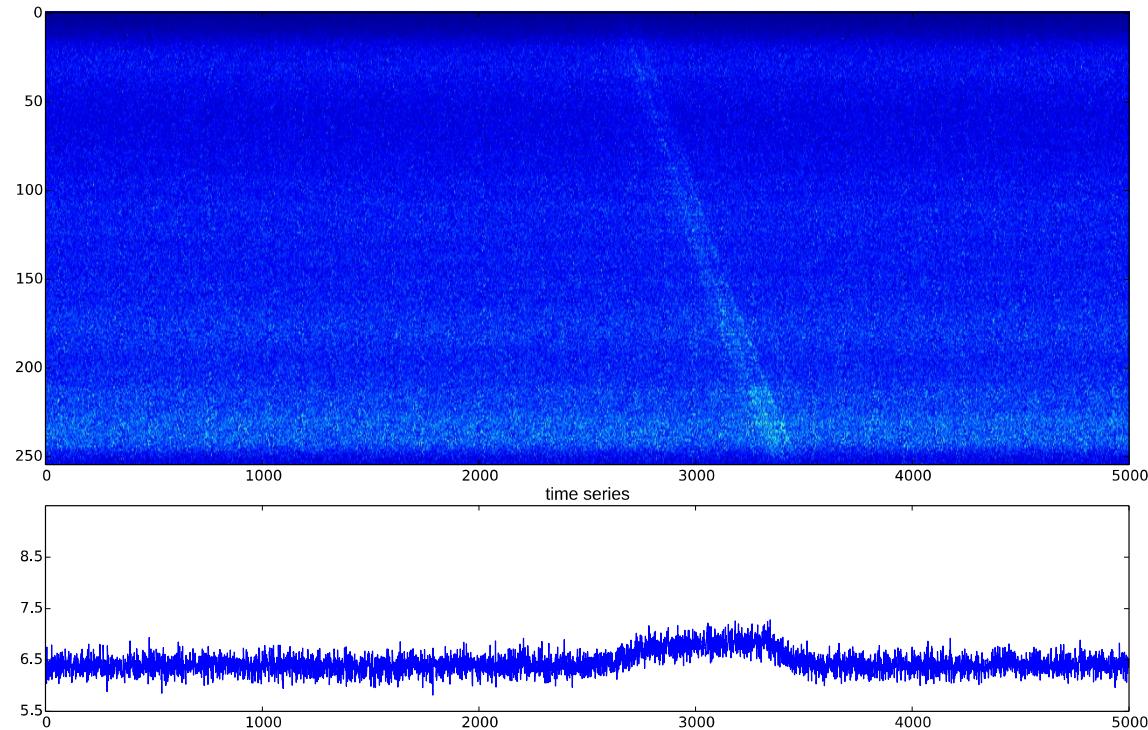
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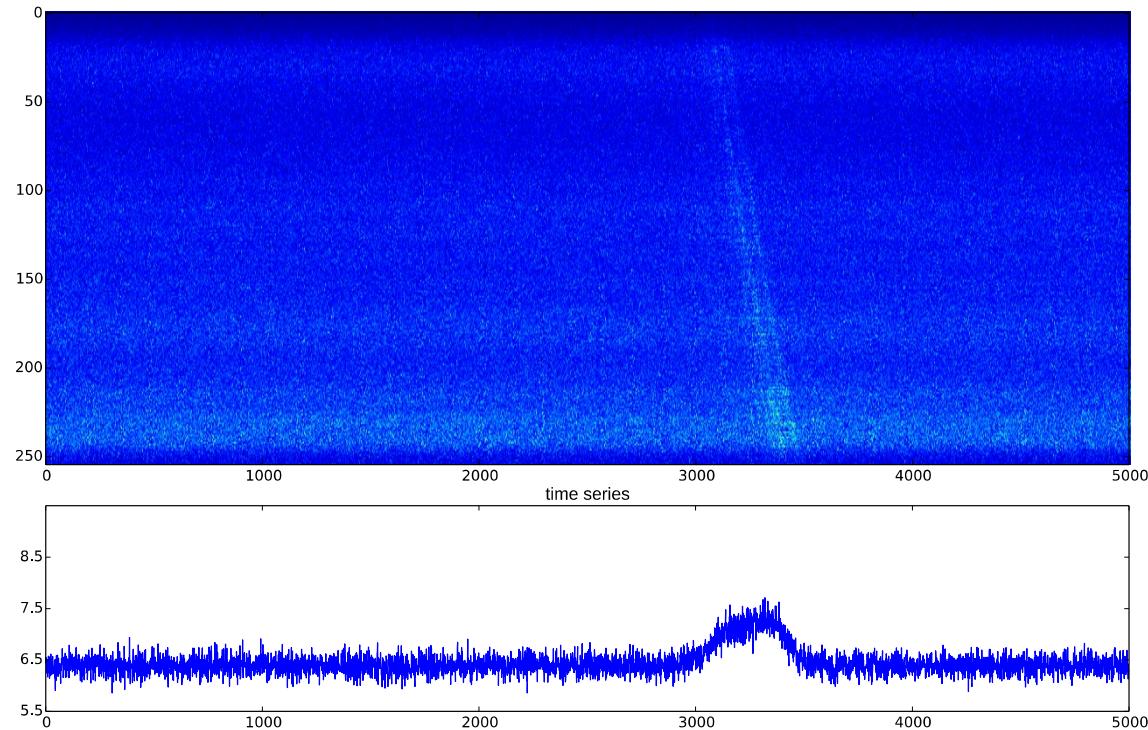
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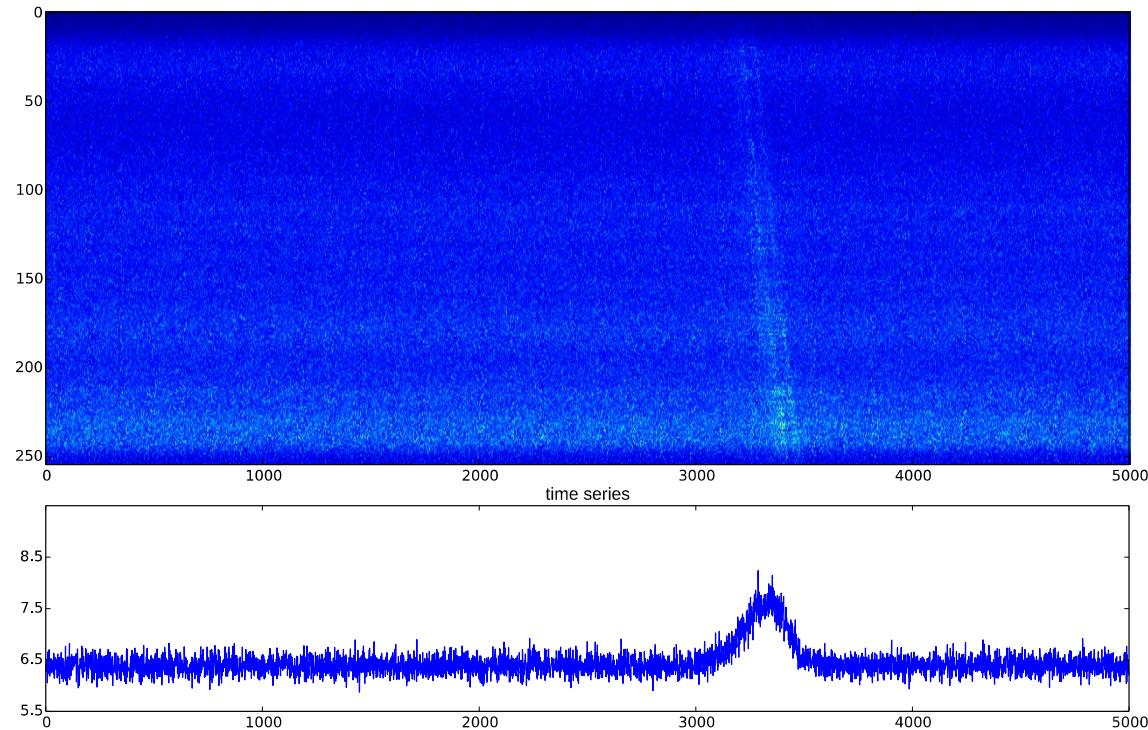
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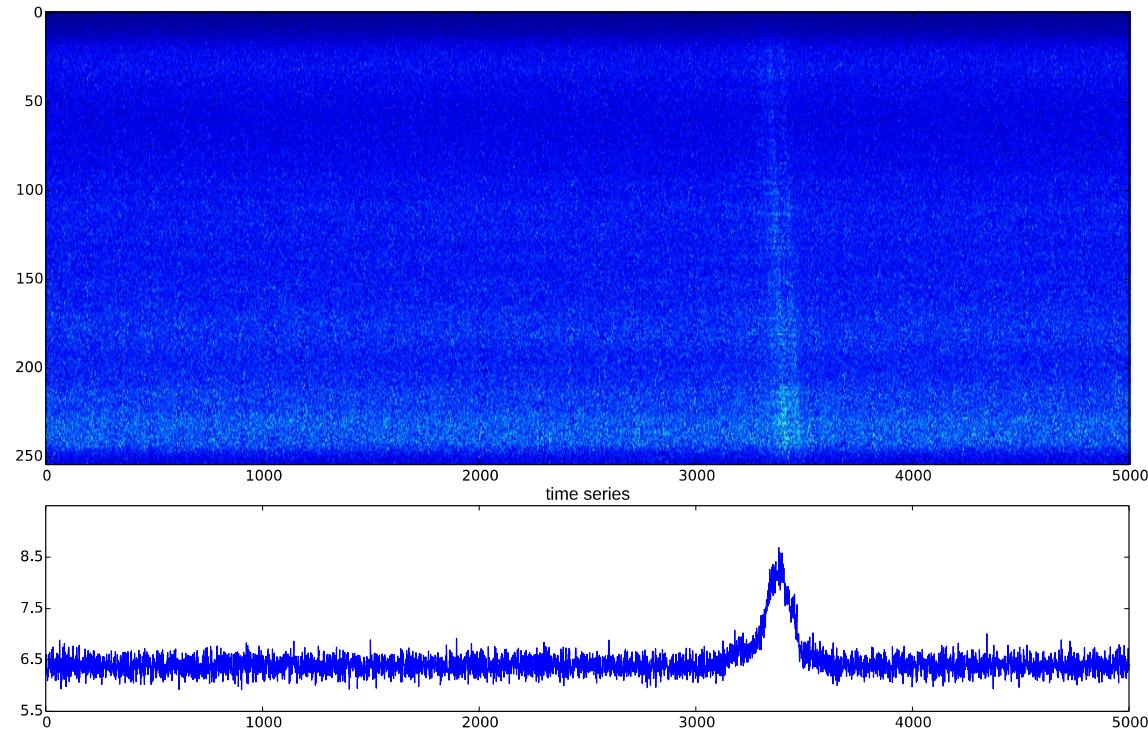
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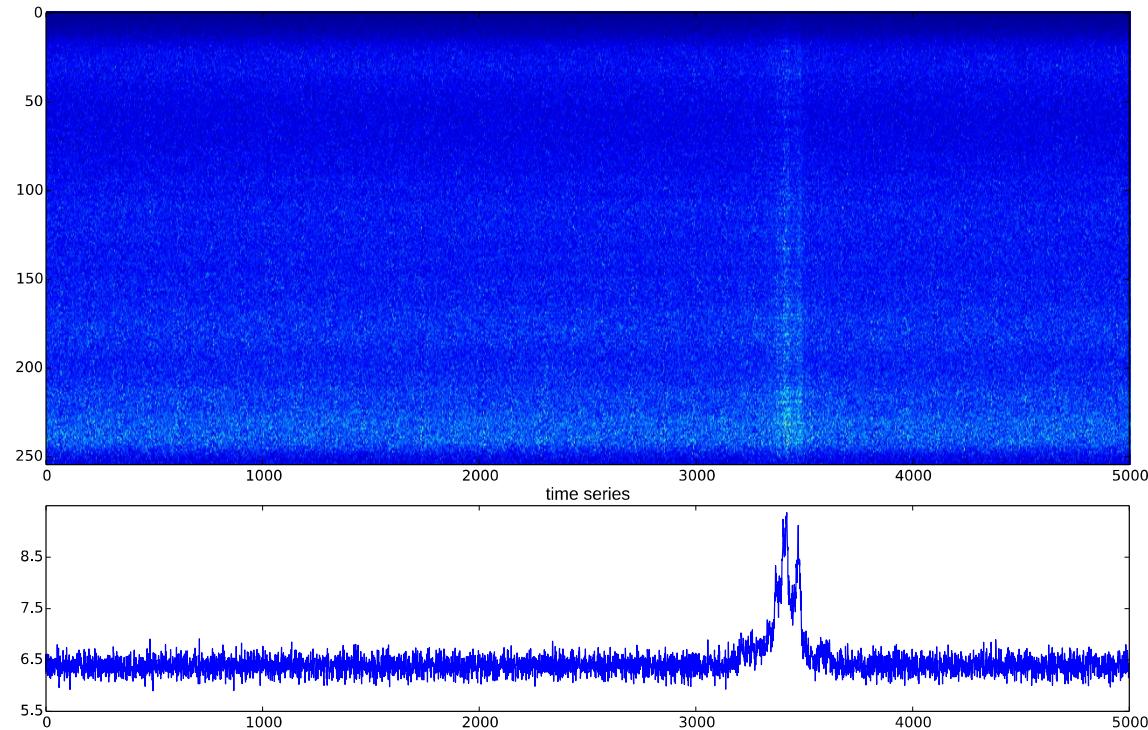
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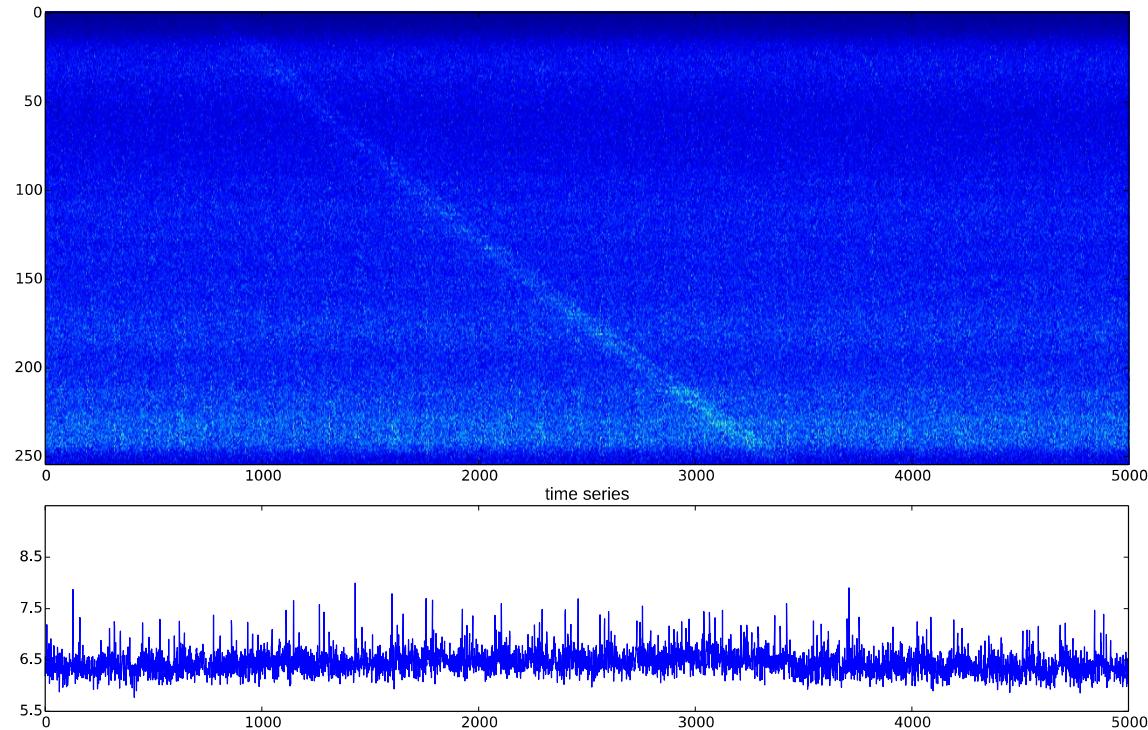
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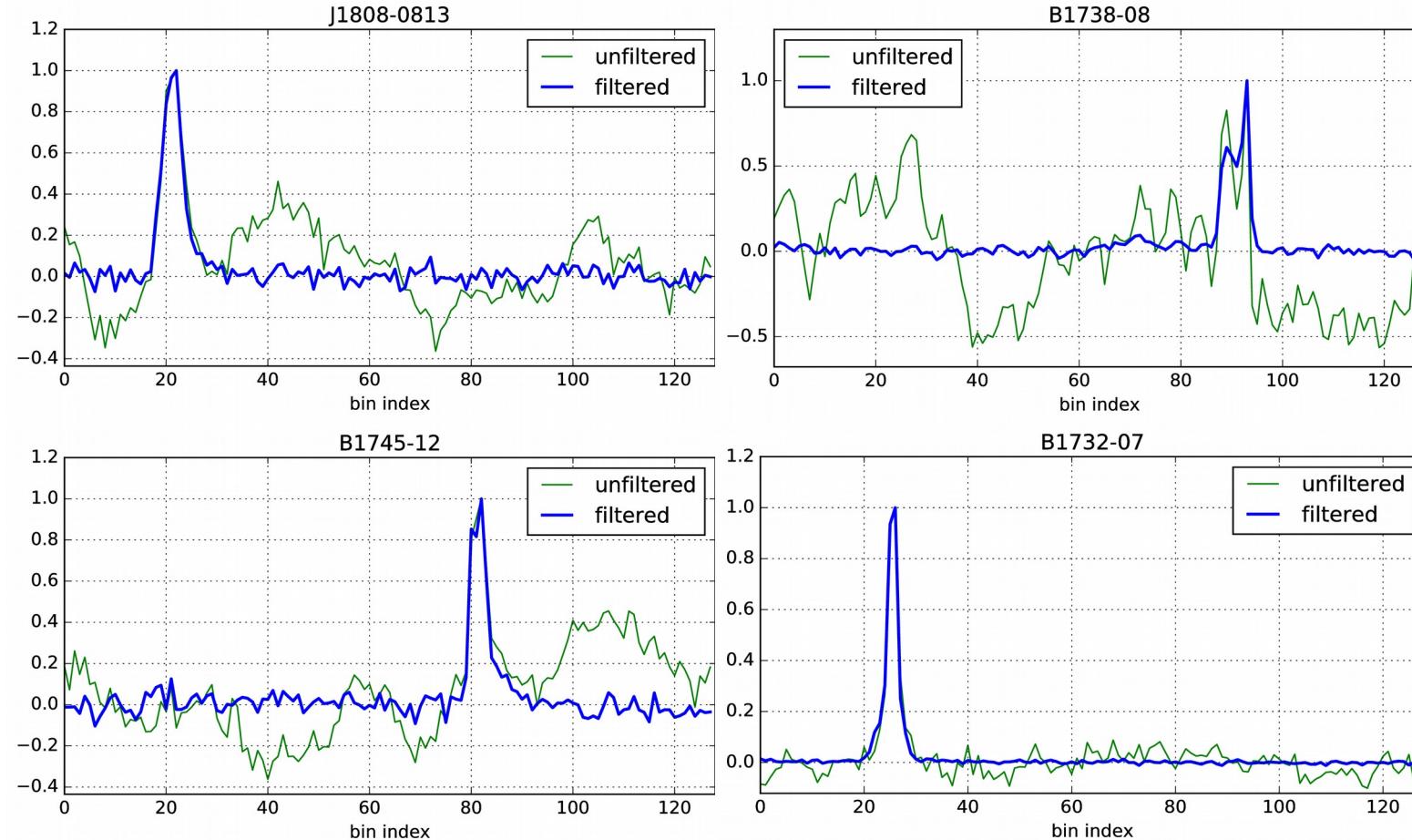
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Impact on Pulsar Data



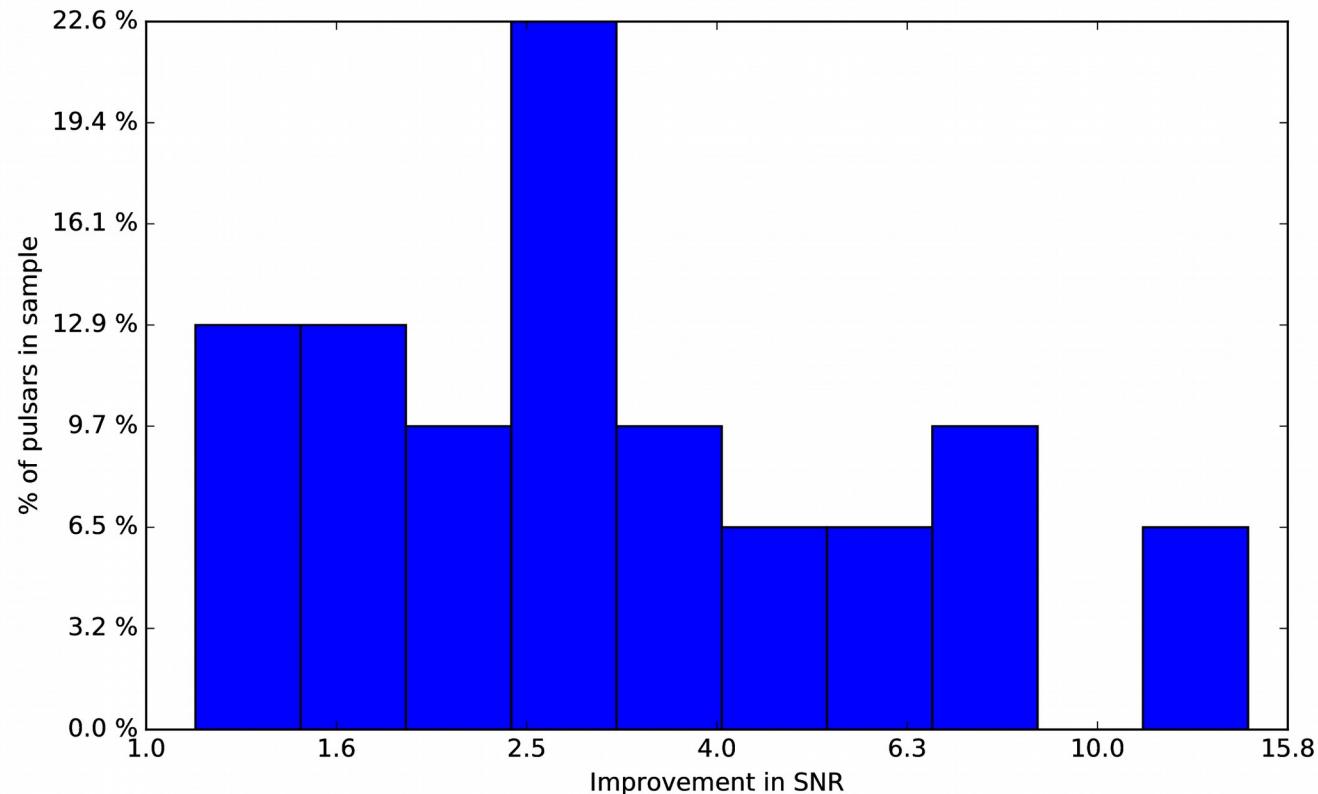
- The tool is regularly being used at GMRT for pre-processing of pulsar data, both for timing and survey.
- On-going pilot survey for pulsars and transients using the uGMRT.
- Results from 32 known pulsars in the field :



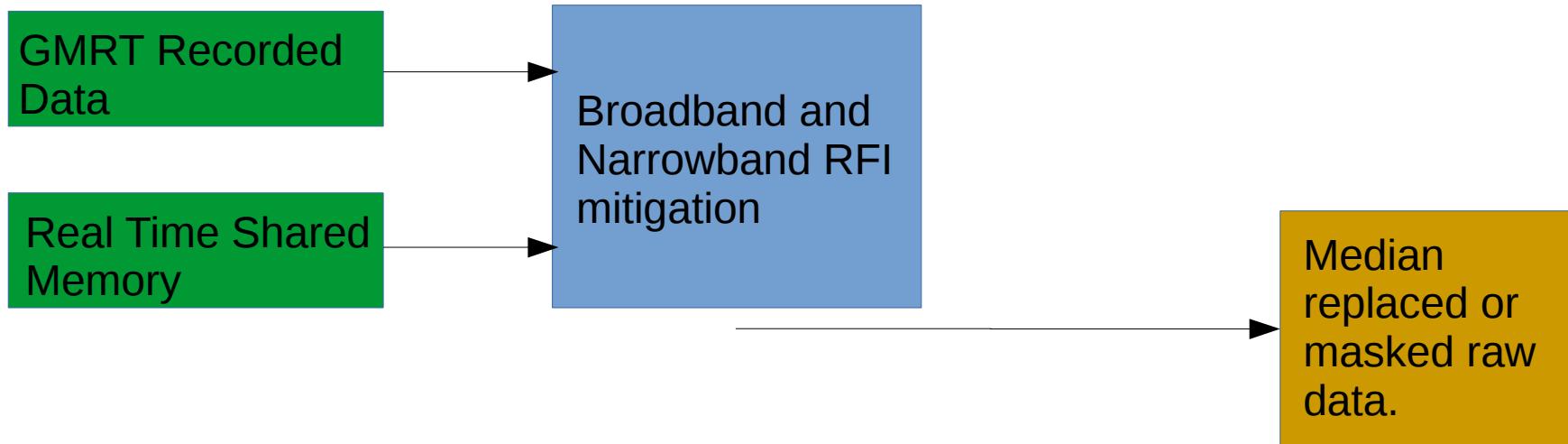
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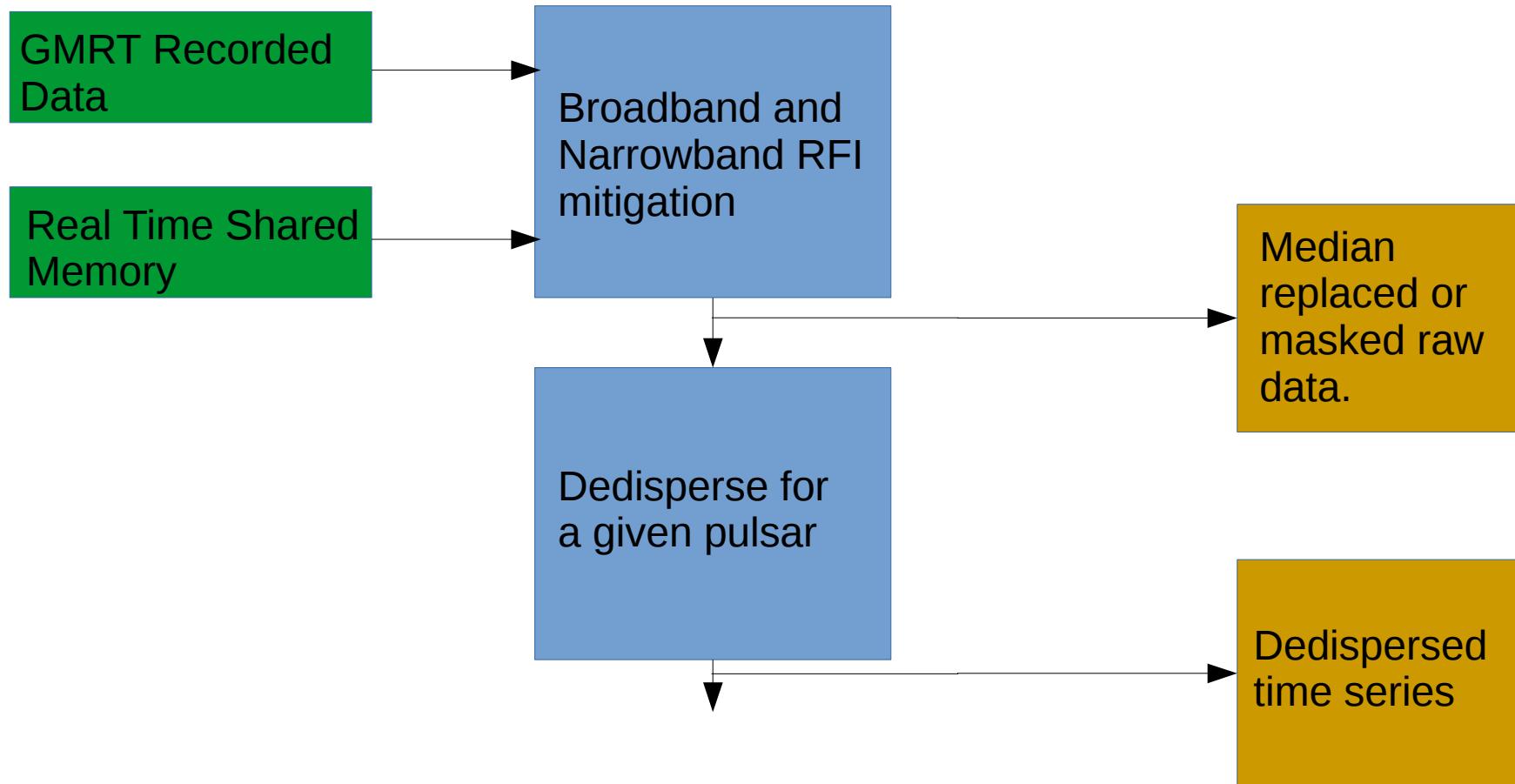
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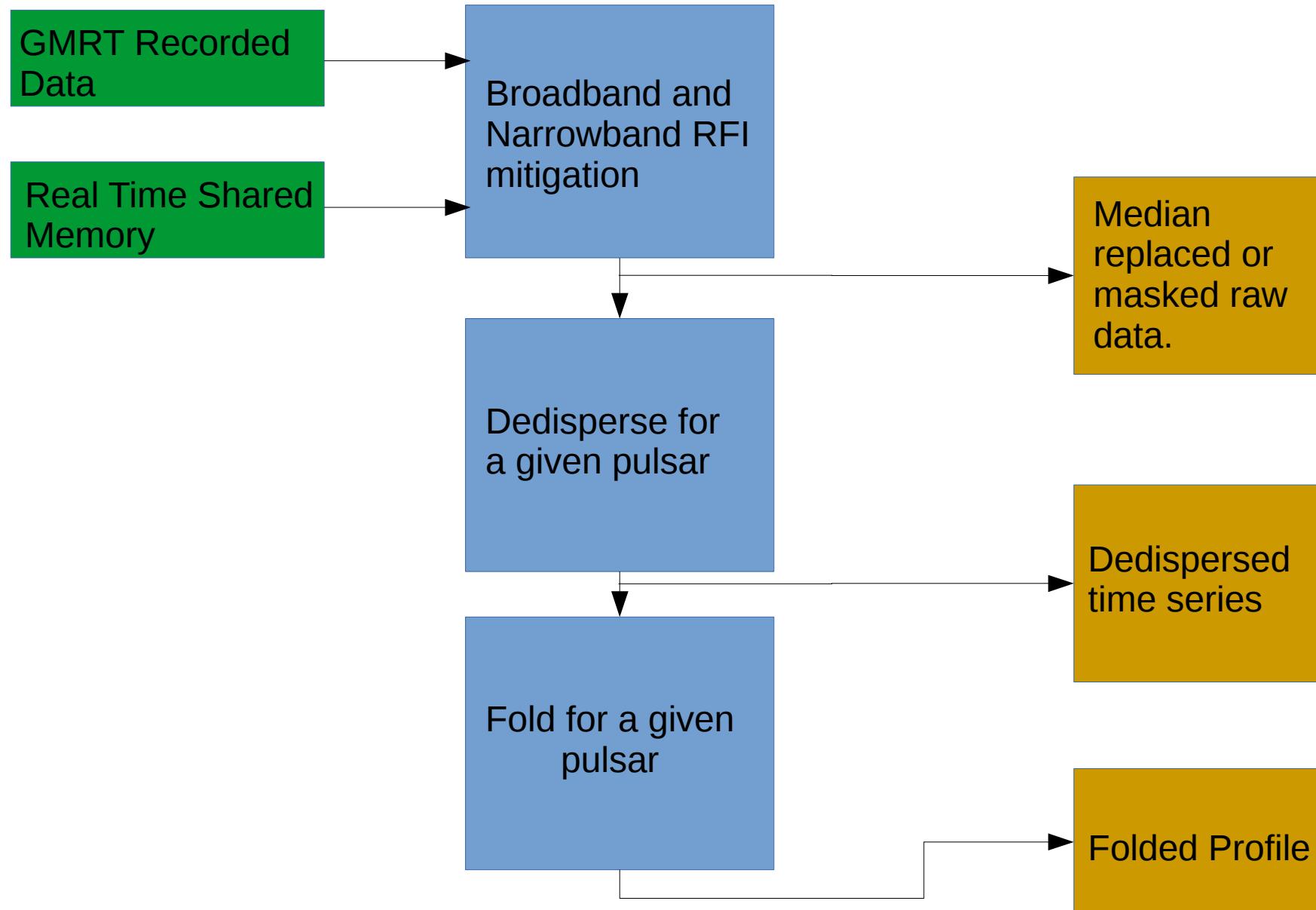
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Need for a Real Time Tool

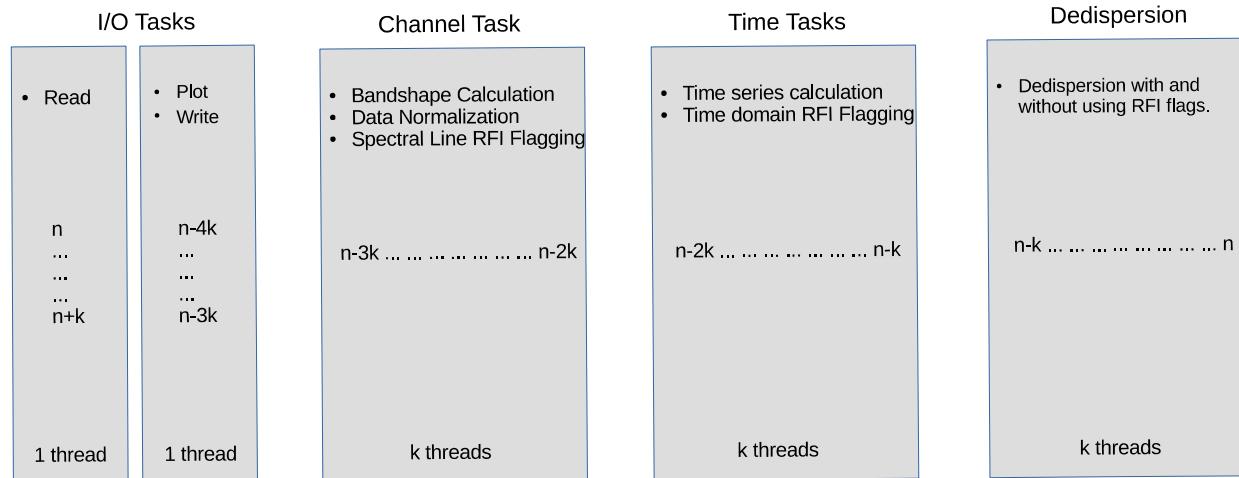


- Nyquist data rate using 200MHz bandwidth of the uGMRT – 400Mb/s
- Can go down to 10 microsecond sampling with 2k spectral channels.
- The resolution can be very useful for pulsar timing studies.
- Writing out full data will be bottlenecked by i/o access.
- Can full time-frequency information be reduced and dumped out in real time?
- **Yes – only if RFI mitigation is efficiently done and only dedispersed time series is dumped out.**

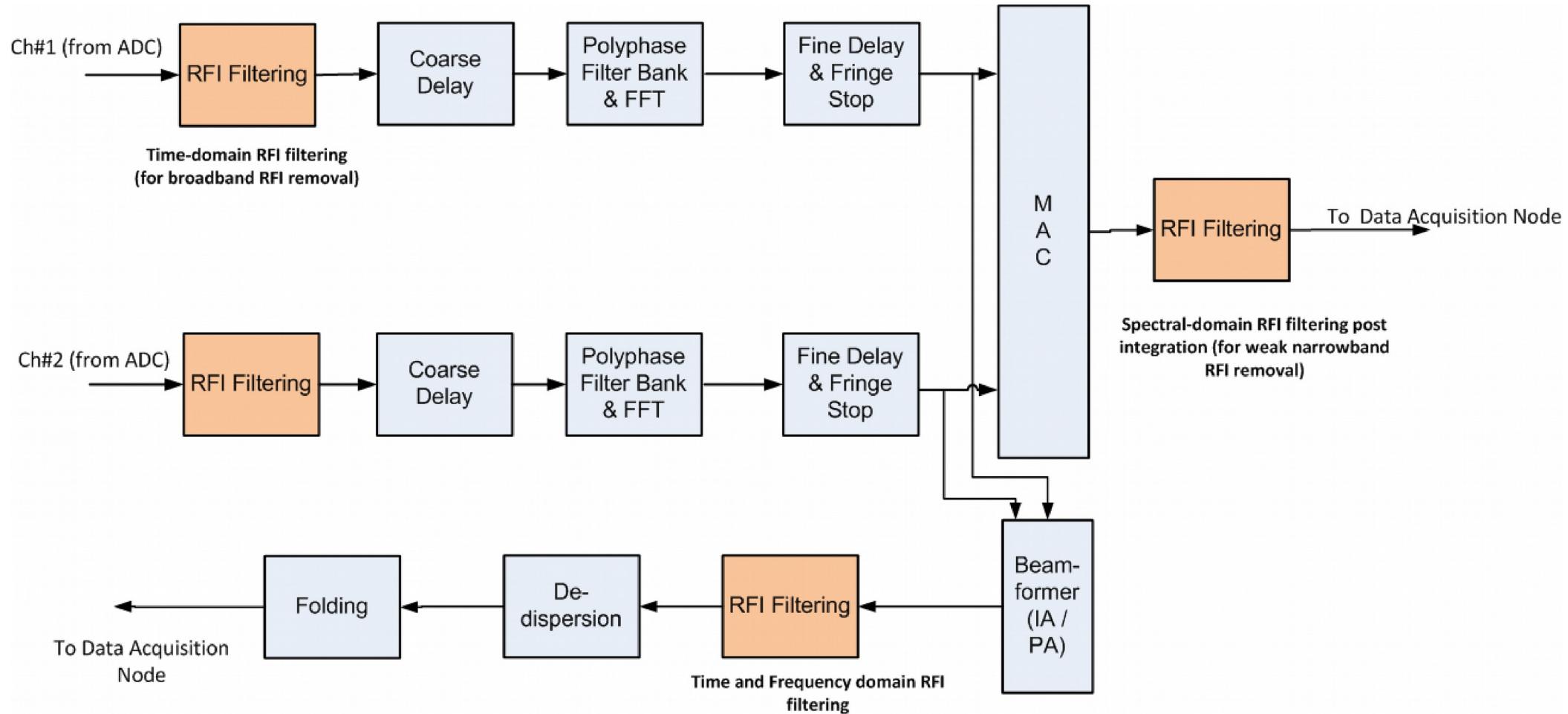
Implementation - The GMRT Pulsar Tool (gptool)



- Written in c++ ; parallelization using OPENMP
- Real time performance:
 - 5 threads by default.
 - Can comfortably process (RFI mitigate -> Dedisperse -> fold -> dump dedisperse series to disk) at a rate of **100 Mb/s**.
 - **40.96 microseconds with 2k spectral channels.**
- Scalability – Use $5 \times N$ threads, effectively handling $100 \times N$ Mb/s.



Real Time RFI Mitigation at GMRT

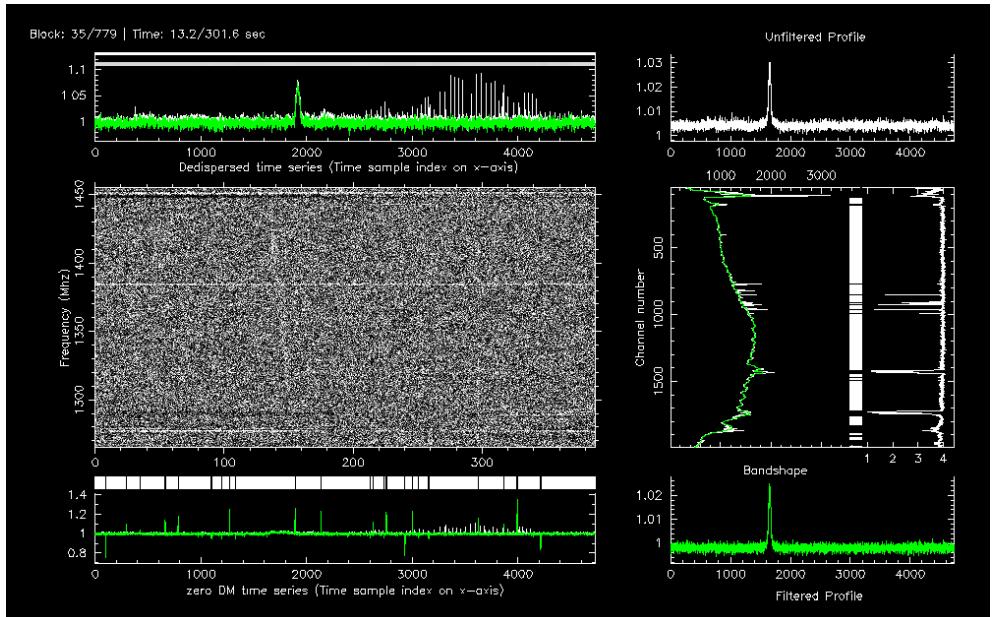


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THE GMRT PULSAR TOOL

- Reliable and proven RFI mitigation strategy with easy to access customization for the user.
- Bundled with basic pulsar processing routines.
- All of it efficiently implemented to run in real time.
- Easy performance scalability – use as much resource as you need!
- Feel free to get in touch for a demo!



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ygupta@ncra.tifr.res.in