

Temporal Evolution of Dynamo-Generated (sub)Stellar Magnetic Fields

Magnetic fields are present in nearly all stars and substellar objects throughout their temporal evolution [hereafter (sub)stellar evolution; Donati & Landstreet, 2009]. (Sub)Stellar magnetic fields come in two flavors: (1) dynamo-generated fields, or magnetic fields that are actively created and sustained by internal dynamical processes, and (2) fossil/remnant fields, which are not actively created or sustained by dynamical processes. Magnetic fields, especially dynamo-generated fields, influence a number of physical processes throughout a star's lifetime, such as a star's rotational evolution, and they may ultimately help govern the physical structure of (sub)stellar interiors. However, investigations studying the role magnetic fields play in regulating physical phenomena regularly neglect the intimate, reciprocal relationship between stellar structure and magnetic field properties.

Purpose and Aims

This proposal aims to understand the relationship between dynamo-generated magnetic fields and the structural evolution of stars and substellar objects. That is, to investigate how the properties of dynamo-generated magnetic fields are affected by (sub)stellar structure over time, as well as how (sub)stellar structure is influenced by the presence of magnetic fields. Furthermore, the proposal seeks to deepen understanding about relationship between dynamo-generated magnetic fields and remnant/fossil magnetic fields that typically exist at later evolutionary stages when dynamo processes subside.

The research program will rely on theoretical modeling of stellar magnetic dynamos simultaneously with theoretical modeling of stellar structure and evolution. Specifically, by coupling a sophisticated one-dimensional stellar evolution model, designed to include effects of magnetic fields on stellar structure, with detailed two- and three-dimensional mean-field magnetic dynamo models. While theoretically driven, the research program will rely on a continuous feedback cycle with cutting edge ground- and space-based observations to assess the validity of theoretical predictions.

Key questions this proposal seeks to answer include:

- How does the topology of a large-scale (sub)stellar magnetic field evolve with time and what are the important characteristics of a star's structure that govern it?
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By answering these questions, the proposal will help address several fundamental...in stellar astrophysics.

- What role do magnetic fields play in governing early contraction rates of young (sub)stellar objects, and subsequently, how do magnetic fields influence masses and ages derived for these young objects?
- How do magnetic activity cycles (analogous to the 11 year solar cycle) vary throughout a star's life?
- How

Survey of the Field

Magnetic fields are a ubiquitous feature of stars and (sub)stellar objects across the Hertzsprung-Russell (HR) diagram. Nearly every class of star is observed to possess a magnetic field or is known to have possessed a magnetic field at some point during its lifetime [Donati & Landstreet, 2009]. Broadly speaking, stellar magnetic fields fall into two categories: (1) dynamo-generated field and (2) fossil/remnant fields. The type of magnetic field possessed by a given star is directly related to its position in the HR diagram. Figure ?? shows an HR diagram and highlights the type of magnetic field possessed by stars in each region. Magnetic fields are pervasive in stars across the HR diagram. Support with observational evidence for magnetic field on the surface of stars at different SpTs from pre-MS through RGB. Evidence for core magnetic fields.

The two observed categories of stellar magnetic fields are thought to be indicative of different mechanisms governing the evolution of the field [Donati & Landstreet, 2009]. Magnetic fields in cool low-mass stars, with their weak large-scale and strong small-scale magnetic field components, are created and sustained through dynamical processes resulting from the interaction between azimuthal rotation and predominantly radial convective flows in their outer layers—a so called “hydromagnetic dynamo” or simply “dynamo” [Parker, 1955]. Dynamo mechanisms Fossil fields are topologically stable magnetic fields that exist predominantly in regions of stars w... In contrast, dynamo-generated magnetic fields are created and sustained through an active magnetic dynamo process that continually...

Note that fossil fields beget dynamo fields, which often beget remnant fields.

General description of the physics as motivated by seminal work (Babcock, Leighton, Cowling, Parker). Supported by detailed (R)MHD simulations. Main results are that the interplay between convection and rotation is critical. Shear vs turbulence. Address virtues and limitations. Mean field simulations.

Dependence on stellar properties; strength and topology. Convection zone depth, resistivity, viscosity, density stratification. Typically in the realm of classical 1D stellar evolution models. Simulations either prescribe the information or motivate it by stellar models.

Dynamically important. They regulate a wide variety of phenomena. Chief among these include driving magnetized stellar winds that ultimately regulate angular momentum evolution and rotational spin-down (Skumanich relation), emission of high energy radiation and eruptive phenomena during activity cycles (important for planetary environments), and they govern the physical structure of stellar interiors. Structure of stars affected by B-field as revealed through 1D models. Has consequences for the development and evolution of magnetic fields. Discuss previous work, key results, highlight how it adds ambiguity to observations. Address virtues and limitations.

In all cases, critical characteristic of magnetic field is the field strength and topology. Both regulate phenomena in different ways. Sometimes it's not entirely clear. Dependence of strength and topology on stellar properties implies dependence on stellar evolutionary state.

Time dependence!! Temporal evolution of dynamo-generated magnetic fields.

Project Description

Describe coupling of stellar evolution and mean-field dynamo theory.

P1. Dynamo-Generated Magnetic Fields throughout Stellar Evolution

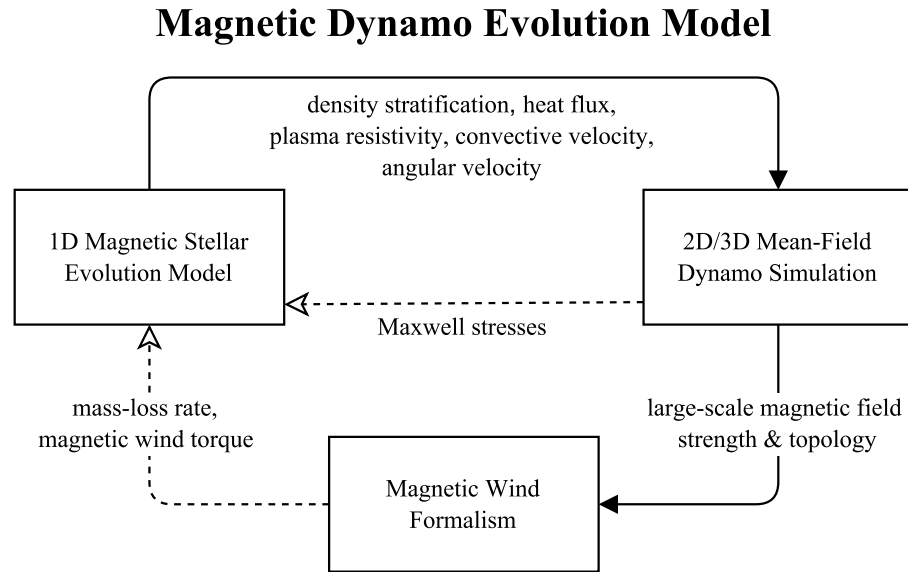


Figure 1: Flow of information in a magnetic dynamo evolution model. A 1D magnetic stellar evolution model [DMESTAR; Feiden & Chaboyer, 2012] forms the system’s backbone, passing interior structure parameters to a 2D/3D mean-field magnetic dynamo simulation [PENCIL code; Brandenburg & Dobler, 2002]. Maxwell stresses computed as a function of radius in the dynamo simulation are passed back to the 1D stellar model while the large-scale magnetic field strength and topology are used to compute magnetic wind breaking properties [e.g., Réville et al., 2015] for the 1D stellar evolution model. The system is iterated in time and the cycle begins anew.

This project is subdivided into three sub-projects denoted SP1, SP2, and SP3. Each sub-project focuses on a particular stage of stellar evolution from the pre-main-sequence through the first ascent red giant branch. The majority of the project’s focus will be on SP1: the early evolution of dynamo generated magnetic fields. Although relatively brief compared to full evolutionary timescales, pre-main-sequence evolution dictates... and has far reaching implications for... Continuing evolution beyond the pre-MS comes at a very small computational cost. Evolution along the main-sequence can be modeled with very little computational time. Recently, magnetic fields have been proposed to explain several observational phenomena revealed by the *Kepler* Space Telescope [Fuller et al., 2015; van Saders et al., 2016]. The code we propose to develop provides an opportunity to study the physics underlying the observed phenomena. Therefore, we will allow stars and their dynamo-generated magnetic fields to evolve along the main-sequence and up the red giant branch.

SP1. Pre-Main-Sequence Evolution

Accurate absolute stellar ages are perhaps the most sought after of astrophysical quantities. This is especially true for identifiably young systems, where absolute ages are our only constraints on time-dependent astrophysical phenomena, such as timescales for dispersal of primordial gaseous circumstellar disks [??] and extrasolar giant planet formation [?]. Furthermore, stellar ages are critical for estimating masses of brown dwarfs and directly imaged giant planets [?]. Despite their importance, accurate ages for young stars and stellar populations remain elusive. This is largely due to a necessary reliance on predictions from stellar evolution models, which are beset with problems at ages younger than 250 Myr [Stassun et al., 2014].

Most concerning is that stellar models predict ages that are strongly dependent on effective temperature and on the stellar properties used in the analysis. HR diagram fitting of stars with $T_{\text{eff}} \gtrsim 5500$ K yield ages that are a factor of two older than the cooler stellar population [Herczeg & Hillenbrand, 2015]. For example, ages in the Upper Scorpius OB association vary by a factor of 2 – 3 for different effective temperature regimes (see Figure 1). Stellar model ages also depend on whether one compares observations to models in the HR or mass-radius diagram. Ages for the same star in Upper Scorpius differ by a factor of two between the HR and mass-radius diagrams [?]. These results are consistent across model sets [Herczeg & Hillenbrand, 2015], indicating problems are endemic to young stellar models and that models are missing physics. A factor of two error in stellar ages means directly imaged young brown dwarf and giant planet masses may be systematically underestimated by 35%.

D’Antona et al. [2000] presented possible hints of the missing physics. They showed that inclusion of rudimentary magnetic inhibition of convection provided better agreement between the masses and HR diagram locations for the T-Tauri components of the eclipsing binary RX J0529.4+0041. They also noted that magnetic inhibition of convection reduced lithium depletion. This trend was later confirmed and leveraged to reconcile the HR diagram and lithium depletion boundary age for the β -Pictoris moving group [MacDonald & Mullan, 2010; Malo et al., 2014]. Notably, Malo et al. [2014] further demonstrated that models including magnetic fields could potentially mitigate age dependencies on spectral type.

Viability of magnetic fields as “the missing physics” in stellar evolution models is best evidenced by *Kepler*/K2 observations of eclipsing binary systems in Upper Scorpius (see Figure 2). Magnetic inhibition of convection appears to provide the missing link needed to relieve spectral type dependencies in stellar ages in Upper Scorpius (Feiden 2016). **Critically, magnetic stellar models naturally reproduce the slope of the mass-radius relationship at the same age inferred from the HR diagram, a feat standard models cannot accomplish.**

SP2. Main-Sequence Evolution

By the time stars reach the hydrogen burning main sequence, the relative age difference predicted by standard and magnetic stellar evolution models is quite negligible. Nevertheless, magnetic fields can continue to play an important role in stellar main sequence evolution. This depends largely on how fast the star is rotating, which in turn, is governed by the topology of the large-scale magnetic field and the efficiency of magnetic wind breaking mechanisms.

As stars evolve along the main sequence, magnetic winds exert a torque on the star, carrying away angular momentum gradually slowing the star’s rotation rate—so called “magnetic wind breaking.” Throughout most of a star’s main sequence lifetime, its rotational velocity decays and is proportional to the inverse square root of its age, $v_{\text{rot}}(t) \propto t^{-1/2}$ [Skumanich, 1972]; a direct consequence of magnetic wind breaking [Durney, 1972]. However, recent evidence shows that magnetic wind breaking is effectively halted during the late stages of a star’s main sequence lifetime, leading to older main sequence stars rotating more rapidly than expected [van Saders et al., 2016].

This sub-project will explore potential physical mechanisms underlying the abrupt weakening of magnetic wind breaking and will seek to correlate the physical mechanisms with properties of stellar interior structure. Multiple hypotheses have been proposed: (i) a sudden decrease in the large-scale magnetic field strength [Durney & Latour, 1977], (ii) a change in the large-scale magnetic field topology [Réville et al., 2015], and (iii) changes in the distribution of

magnetically active regions at the stellar surface [Cohen et al., 2009].

The coupled stellar dynamo-evolution model will be uniquely suited to explore hypotheses (i) and (ii).

Exploration of option (iii) is difficult. Limited spatial resolution makes it difficult to discern small-scale features such as magnetic active regions. In addition, global mean-field dynamo simulations do not explicitly account for Lorentz-force feedback on convective flows, meaning suppression of convection necessary to create starspot groups is not possible. Our approach implicitly assumes that magnetic inhibition of convection is a global phenomenon and occurs over thermal timescales allowing us to treat it in the stellar structure calculation [Feiden & Chaboyer, 2012].

SP3. Post-Main-Sequence Evolution

Depletion of hydrogen in stellar cores eventually reaches a level... Stars begin to traverse across the HR diagram along the sub-giant branch and up the first-ascent red giant branch once hydrogen shell burning ceases. Stars more massive than approximately 1.2 Msun possess a convective core along the hydrogen burning main sequence that is able to sustain a dynamo-generated magnetic field [e.g., Browning et al., 2004; Featherstone et al., 2009; Augustson et al., 2016]. Upon leaving the main sequence, however, convection in the stellar core rapidly subsides (8 Myr). The fate of the dynamo-generated magnetic fields produced by the convective core during this period of rapid transition off the main-sequence is unexplored.

There is indirect evidence from asteroseismology that remnant magnetic fields exist in the convectively stable cores of in a fraction of red giant branch stars. Most red giant branch stars observed by Kepler pulsate at specific frequencies that are characteristic of a coupling between wave-modes generated in the radiative stable core and modes created in convectively unstable outer envelope. However, a fraction of stars in the Kepler exhibit a very weak, if not non-existent, coupling between these wave modes [Fuller et al., 2015]. The current hypothesis is that strong magnetic fields left over from the dynamo-generation process inhibit the propagation of g-modes in the radiative core, a so-called “magnetic greenhouse effect,” ultimately suppressing the coupling of modes between the core and envelope [Fuller et al., 2015; Cantiello et al., 2016]. However, it is unclear whether rapidly shrinking convective cores are able to produce remnant magnetic fields that are strong enough (strength) or possess a proper topology (predominantly dipolar?) to efficiently suppress mode coupling.

We will study how the strength and topology of dynamo-generated magnetic fields evolve in the cores of intermediate-mass stars leading up to, during, and following their transition off the main sequence. Detailed magnetohydrodynamic simulations show that convective cores can sustain magnetic dynamos and that the resulting magnetic field can affect the properties of convection.

P2. Dynamo-Generated Magnetic Fields in Brown Dwarfs and Giant Planets

Delayed contraction in stars. What about in brown dwarfs and giant planets? Affect their thermal history and thus masses that we infer for directly imaged objects. Impacts formation theories (via mass distribution) and even the status of an object (star vs bd, bd vs planet).

P3. Fundamental Properties of Young Single Stars

Inferring accurate ages for young stars requires that stellar evolution models predict accurate fundamental properties for real stars. *This project focuses on confronting model predictions of young star properties as a function of age using a large number of well-characterized stars*

and young eclipsing binaries observed by Kepler/K2. Previous investigations have found that young stellar models are unable to accurately predict effective temperatures and luminosities for stars at a given mass [e.g., Hillenbrand & White, 2004; Mathieu et al., 2007; Stassun et al., 2014]. These studies have focused on astrometric and eclipsing binaries (EBs), making it difficult to rule out whether binarity is playing a role in driving disagreements. Even wide binaries may be subject to subtle interactions that corrupt agreement with models. Recently, Stassun et al. [2014] showed that, while binarity in and of itself may not be a problem, model errors appear to be correlated with the presence of a tertiary companion. Triplicity among the sample of eclipsing binaries in Stassun et al. [2014] was well above 50%, reducing the sample of potentially reliable comparisons to models. Our understanding about whether young stellar models are accurate at predicting properties of single stars may therefore be biased.

Therefore, we are creating the first large sample of single young benchmark stars, whose fundamental properties will be measured with high precision (better than 5% for young stars). Stellar properties are derived using flux-calibrated spectra obtained by collaborators Mann and Gaidos. The spectra are used to derive bolometric fluxes and effective temperatures, leading to empirical estimates of stellar radii. This technique was proven to be robust for assessing stellar evolution model reliability for main sequence field stars [Mann et al., 2015]. Our current target sample contains 130 young stars in nearby moving groups and young stellar associations with known distances.

Comparison of properties of benchmark stars to stellar models will be performed using Bayesian inference to derive best fit model properties with realistic uncertainties [Mann et al., 2015]. We will then explore how well models reproduce the observed data and whether there are any noticeable correlations with additional observables, such as magnetic activity indicators, inferred age, chemical composition, and the convective mixing length parameter. Since we are targeting stars that are members of known young moving groups and associations, we will check that models yield consistent ages and distances to stars from a common group. If there are inconsistencies, we will attempt to characterize them in terms of observables and inferred stellar model parameters.

P4. Homogeneous Ages for Young Stellar Associations

Deriving homogeneous ages for young stellar associations using available literature data. Provides constraints on a number of interesting astrophysical phenomena. [partially with UT Austin; Stockholm]

Informed by results obtained in projects 1, 2, and 3, we will bring together all of the information to establish ages for young stellar systems that are consistent for different age measures. The age determinations will fold in all necessary ingredients, starting with a base set of standard young stellar models and exploring whether magnetic fields, starspots, or chemical composition are able to provide complete agreement for each system. We'll explore the possibility that the appropriate mix of ingredients from work package 1 and other modified physics from work package 2 lead to consistent ages for all systems or if there are correlations among the mix of ingredients and stellar system parameters like age.

To avoid being entirely dependent on the success of WP01 and WP02, we will begin an initial effort characterizing how magnetic stellar models developed by GAF perform in terms of predicting consistent ages given different age indicators. Notably, consistency between high- and low-mass regions of CMDs and the HRD, as well as the lithium depletion boundary. We have preliminary results using a single eclipsing binary in Upper Sco (Kraus et al. 2015) that

indicates magnetic models may provide favorable consistency between different mass regimes and between different age indicators for the binary alone (see Figure 2). We will therefore try to develop a more complete analysis of Upper Sco, including more archival data, and other nearby YMGs to determine the robustness of the solution. We will also provide our magnetic models to the community to permit additional independent efforts to evaluate the success or failures of these models.

Results of this exercise will be important, whether or not we are able to derive consistent ages or the different stellar systems. If we can derive consistent ages, then that will provide strong support for those derived ages, at least for the moment, which will become a powerful resource for the astronomical community. If we cannot, then we learn that there is yet something missing from models or that our initial hypothesis that low-mass stars are to blame for the systematic age differences observed in young stellar systems is incorrect.

Finally, this sub-project requires production of a grid of pre-main-sequence magnetic stellar models. Since the grid will be produced concurrently with the above studies, no additional time will be required. Local clusters at Uppsala will provide computational resources necessary and the PI is applying for time with SNIC, the Swedish high-performance computing center. All models, data products, and analysis tools will be made publicly available.

Significance

- Ages of Young Stars
 - protoplanetary disk evolution timescales
 - giant planet formation timescales
 - star formation history, including time dependence of IMF
 - properties of directly imaged planets and brown dwarfs (mass distribution → formation)
- Magnetic Field Evolution
 - activity cycles → astrobiology (stretch?)
 - angular momentum transport and stellar rotational evolution
 - accurate stellar properties
 - * ages and masses of single stars
 - * properties of transiting extrasolar planets
- Grid of Stellar Models
 - available to public
 - allows extensive testing of theory
- Relevance for Next-Generation Surveys

Preliminary Results

Feiden (2016, submitted). Identification of magnetic fields as culprit for age discrepancies.

Pilot study for the Sun (and low-mass star?).

Feiden & Christophe (in preparation). Parametrized model of starspots.

Stars characterized for CALYPSO thus far.

Independent Line of Research

- Calibration of Low-Mass Young Stars and their Planetary Systems with Observations (CALYPSO): second avenue of research is characterizing planets around young stars. Feiden (projectledare) provides a comprehensive analysis of host stars using stellar models coupled with Bayesian inference to give stellar mass, age, and distance estimates with realistic uncertainties.
- 4MOST Internal Working Group 7 – Galactic Analysis Pipeline – providing stellar structure and atmosphere code for self-consistent determination of atmospheric parameters, mass, and ages for over 1 million stars.
- Zodiacal Exoplanets in Time (ZEIT).

International and National Collaborations

Confirmed Star formation group at UT Austin (Kraus, Rizzuto). Fundamental properties of young stars using B- through M-type EBs in young associations with Kepler/K2. Magnetic fields of young stars with IGRINS.

Confirmed CALYPSO: Calibration of Low-mass Young stars and their Planetary Systems with Observations (Mann [UT Austin], Gaidos [Hawaii], Ansdell [Hawaii]). Fundamental properties of bona fide members of nearby young moving groups using "spectro-bolometric techniques."

Pending MDI/ZDI research group at ESO (Hussain, Lavail). Magnetic field topologies for young T-Tauri stars that can be compared against predictions from models in P1 with age information from P3.

Pending MDI/ZDI research group at Uppsala (Kochukhov, Rosén). Magnetic field topologies for young solar analogs in young stellar associations.

Confirmed Open Cluster collaboration (Silvester, Landstreet). Magnetic field topologies for intermediate mass stars in several young stellar associations.

Other Grants

References

- Augustson K. C., Brun A. S., Toomre J., 2016, ArXiv e-prints
- Brandenburg A., Dobler W., 2002, Comput. Phys. Commun., 147, 471
- Browning M. K., Brun A. S., Toomre J., 2004, ApJ, 601, 512
- Cantiello M., Fuller J., Bildsten L., 2016, ArXiv e-prints
- Cohen O., Drake J. J., Kashyap V. L., Gombosi T. I., 2009, Astrophys. J., 699, 1501
- D’Antona F., Ventura P., Mazzitelli I., 2000, ApJ, 543, L77
- Donati J.-F., Landstreet J. D., 2009, ARA&A, 47, 333
- Durney B., 1972, NASA Spec. Publ., 308, 282
- Durney B. R., Latour J., 1977, Geophys. Astrophys. Fluid Dyn., 9, 241
- Featherstone N. A., Browning M. K., Brun A. S., Toomre J., 2009, ApJ, 705, 1000
- Feiden G. A., Chaboyer B., 2012, ApJ, 761, 30
- Fuller J., Cantiello M., Stello D., Garcia R. A., Bildsten L., 2015, Science, 350, 423
- Hecceg G. J., Hillenbrand L. A., 2015, arXiv: 1505.06518
- Hillenbrand L. A., White R. J., 2004, ApJ, 604, 741
- MacDonald J., Mullan D. J., 2010, ApJ, 723, 1599
- Malo L., Doyon R., Feiden G. A., Albert L., Lafrenière D., Artigau É., Gagné J., Riedel A., 2014, ApJ, 792, 37
- Mann A. W., Feiden G. A., Gaidos E., Boyajian T., von Braun K., 2015, ApJ, 804, 64
- Mathieu R. D., Baraffe I., Simon M., Stassun K. G., White R., 2007, in Protostars & Planets V, pp. 411–425
- Parker E. N., 1955, ApJ, 122, 293
- Réville V., Brun A. S., Matt S. P., Strugarek A., Pinto R. F., 2015, Astrophys. J., 798
- Skumanich A., 1972, ApJ, 171, 565
- Stassun K. G., Feiden G. A., Torres G., 2014, New Astronomy Reviews, 60, 1
- van Saders J. L., Ceillier T., Metcalfe T. S., Aguirre V. S., Pinsonneault M. H., García R. A., Mathur S., Davies G. R., 2016, Nature, 529, 181