

Empirical Constraints on the Nucleosynthesis of Nitrogen

James W. Johnson,^{1,2}★ David H. Weinberg,^{1,2,3} Fiorenzo Vincenzo,^{1,2,4} Jonathan C. Bird,⁵ and Emily J. Griffith^{1,2}

¹ *Department of Astronomy, The Ohio State University, 140 W. 18th Ave., Columbus, OH, 43210, USA*

² *Center for Cosmology and Astroparticle Physics (CCAPP), The Ohio State University, 191 W. Woodruff Ave., Columbus, OH, 43210, USA*

³ *Institute for Advanced Study, 1 Einstein Dr., Princeton, NJ, 08540, USA*

⁴ *E.A. Milne Centre for Astrophysics, University of Hull, Cottingham Rd, Kingston upon Hull, HU6 7RX, United Kingdom*

⁵ *Department of Physics & Astronomy, Vanderbilt University, 2301 Vanderbilt Place, Nashville, TN, 37235, USA*

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ABSTRACT

We derive empirical constraints on the nucleosynthetic yields of nitrogen by incorporating N enrichment into our previously developed and empirically tuned multi-zone galactic chemical evolution model. We adopt a metallicity-independent (“primary”) N yield from massive stars and a metallicity-dependent (“secondary”) N yield from AGB stars. In our model, galactic radial zones do not evolve along the observed [N/O]-[O/H] relation, but first increase in [O/H] at roughly constant [N/O], then move upward in [N/O] via secondary N production. By $t \approx 5$ Gyr, the model approaches an equilibrium [N/O]-[O/H] relation, which traces the radial oxygen gradient. We find good agreement with the [N/O]-[O/H] trend observed in extra-galactic systems if we adopt an IMF-averaged massive star yield $y_{\text{N}}^{\text{CC}} = 3.6 \times 10^{-4}$, consistent with predictions for rapidly rotating progenitors, and a fractional AGB yield that is linear in mass and metallicity $y_{\text{N}}^{\text{AGB}} = (9 \times 10^{-4})(M_*/M_{\odot})(Z_*/Z_{\odot})$. This model reproduces the [N/O]-[O/H] relation found for Milky Way stars in the APOGEE survey, and it reproduces (though imperfectly) the trends of stellar [N/O] with age and [O/Fe]. The metallicity-dependent yield plays the dominant role in shaping the gas-phase [N/O]-[O/H] relation, but the AGB time-delay is required to match the APOGEE stellar age and [O/Fe] trends. If we add $\sim 40\%$ oscillations to the star formation rate, the model reproduces the scatter in gas-phase [N/O] vs. [O/H] observed in external galaxies by MaNGA. We also construct models using published AGB yields and examine their empirical successes and shortcomings. For all AGB yields we consider, simple stellar populations release half their N after only ~ 250 Myr.

Key words: methods: numerical – galaxies: abundances – galaxies: evolution – galaxies: star formation – galaxies: stellar content

1 INTRODUCTION

From a nucleosynthesis perspective, N is a unique element. Along with C and He, it is one of only three elements lighter than iron peak nuclei thought to owe a significant portion of its abundance to asymptotic giant branch (AGB) stars (e.g. [Johnson 2019](#)). N is also the primary by-product of the CNO cycle, a cyclic nuclear reaction that catalyses the conversion of H into He in stars more massive than the sun. However, uncertainties surrounding the nucleosynthetic yields of N make it difficult to model its abundances accurately. Here we take an empirical approach to constrain N yields by using state-of-the-art galactic chemical evolution (GCE) models to assess which functional forms describing the yield can reproduce recent observational data for gas phase abundances and trends in Milky Way disc stars found by [Vincenzo et al. \(2021\)](#).

Observationally, N abundances in external galaxies are generally measured in the gas phase and are used as a metallicity indicator because of their strong correlation with O abundances. In Fig. 1, we present a compilation of such measurements along with data from the Milky Way:

1. HII regions in the first six CHAOS¹ galaxies: NGC 3184, NGC 628, NGC 5194, NGC 5457, M101, and NGC 2403 ([Berg et al. 2020](#); [Skillman et al. 2020](#); [Rogers et al. 2021](#)).
2. HII regions in nearby NGC spirals ([Pilyugin, Vílchez & Thuan 2010](#)).
3. HII regions in blue, diffuse star forming dwarf galaxies ([Berg et al. 2012](#); [Izotov, Thuan & Guseva 2012](#); [James et al. 2015](#)).
4. Local stars and HII regions ([Dopita et al. 2016](#)).
5. Galactic and extragalactic HII regions ([Henry, Edmunds & Köppen 2000](#)).
6. Star-forming regions in 550 nearby galaxies in the MaNGA IFU² survey ([Belfiore et al. 2017](#); [Bundy et al. 2015](#)).

¹ CHAOS: Chemical Abundances Of Spirals ([Berg et al. 2015](#))

² MaNGA: Mapping Nearby Galaxies at Apache Point Observatory ([Bundy et al. 2015](#)). IFU: Integral Field Unit.

Despite intrinsic scatter and some systematic variation in how the abundances are determined, this $[N/O]-[O/H]$ ³ relation is found to be similar across a wide range of astrophysical environments. Furthermore, recent arguments from both theoretical ([Vincenzo & Kobayashi 2018](#)) and observational perspectives ([Hayden-Pawson et al. 2021](#)) suggest that this relation is largely redshift-invariant. Previous studies have interpreted this consistency as an indication that the relation is nucleosynthetic in origin, reflective of a “primary” yield that does not depend on a star’s initial metal content and a “secondary” yield that does ([Vila-Costas & Edmunds 1993](#); [van Zee, Salzer & Haynes 1998](#); [Henry & Worthey 1999](#); [Pérez-Montero & Contini 2009](#); [Pilyugin, Grebel & Mattsson 2012](#); [Andrews & Martini 2013](#)). Although we have highlighted star forming galaxies in Fig. 1, N abundances are also easily measured in massive ellipticals (see, e.g., [Schiavon 2010](#), [Conroy 2013](#), and [Conroy, Graves & van Dokkum 2014](#) for observational references), allowing it to potentially bridge the gap between the physical processes affecting galaxies of different morphologies.

The challenge in interpreting N abundances is that accurate nucleosynthetic yields from various enrichment channels remain elusive. Relative to other light elements, N synthesis is difficult to model because it is sensitive to uncertain details of stellar evolution, such as internal mixing (see discussion in, e.g., [Andrews et al. 2017](#) and in § 2.1 below). In this paper, we constrain N yields empirically by testing the performance of various assumptions within the framework of GCE models. To this end we make use of the multi-zone model for the Milky Way published by [Johnson et al. \(2021\)](#), which treats the Galaxy as a series of concentric rings, describing each one as a conventional one-zone model of chemical evolution (see discussion in § 3). This approach has been employed in the past to compute abun-

³ We follow standard notation where $[X/Y] \equiv \log_{10}(X/Y) - \log_{10}(X/Y)_{\odot}$.

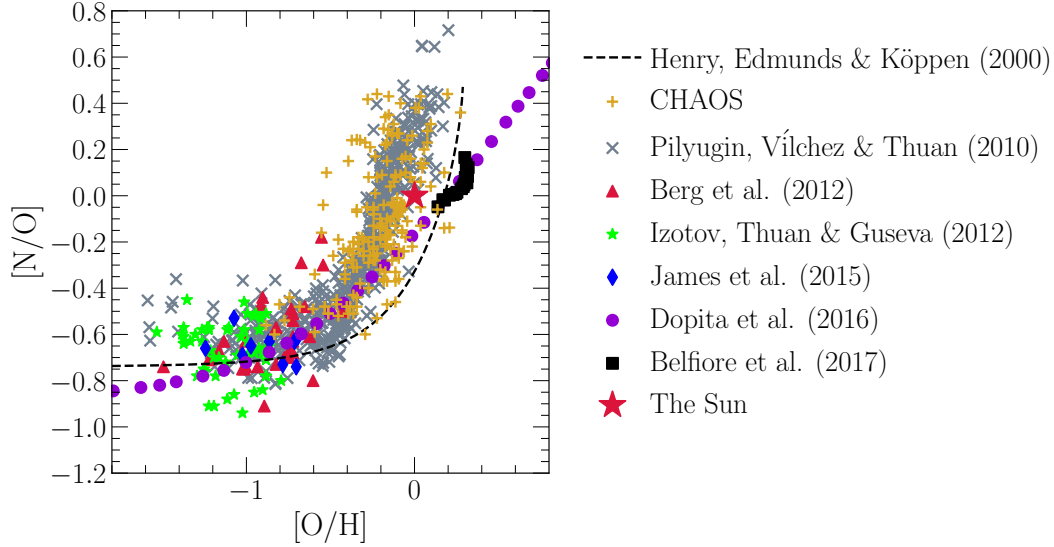


Figure 1. The $[N/O]$ - $[O/H]$ relation as observed in different galactic environments: HII regions from the first six CHAOS galaxies (golden +’s: NGC 3184, NGC 628, NGC 5194, NGC 5457, M101, and NGC 2403; Berg et al. 2020; Skillman et al. 2020; Rogers et al. 2021) and other nearby NGC spiral galaxies (grey X’s; Pilyugin et al. 2010), HII regions in blue diffuse star forming dwarf galaxies (red triangles: Berg et al. 2012; green stars: Izotov et al. 2012; blue diamonds: James et al. 2015), in local stars and HII regions (purple circles: Dopita et al. 2016), and in the MaNGA IFU survey (black squares: Belfiore et al. 2017). The fit to $[N/O]$ as a function of $[O/H]$ in Galactic and extragalactic HII regions by Henry et al. (2000) is shown as a black dashed line. We omit all uncertainties for visual clarity. The Sun, at (0, 0) on this plot by definition, is marked by a large red star.

dances for many Galactic regions simultaneously (Matteucci & Francois 1989; Wyse & Silk 1989; Prantzos & Aubert 1995; Schönrich & Binney 2009; Minchev, Chiappini & Martig 2013, 2014; Minchev et al. 2017; Sharma, Hayden & Bland-Hawthorn 2021). Because of the apparent universality of the $[N/O]$ - $[O/H]$ relation, our results using the Milky Way as a case test should apply to other galaxies as well.

At low metallicity, rotating massive stars play a key role in establishing the observed N abundances (Chiappini, Romano & Matteucci 2003; Chiappini, Matteucci & Ballero 2005; Chiappini et al. 2006; Kobayashi, Karakas & Umeda 2011; Prantzos et al. 2018; Grisoni, Matteucci & Romano 2021). Rotation plays a pivotal role in stellar evolution, inducing effects such as shear mixing, meridional circulation, and horizontal turbulence (Zahn 1992; Maeder & Zahn 1998; Lagarde et al. 2012). These effects carry internally produced C and O nuclei into the H-burning shell where they can be processed into ^{14}N via the CNO cycle (Heger & Woosley 2010; Frischknecht et al. 2016; Andrews et al. 2017). Metal-poor stars spin faster and are more compact (Maeder, Grebel & Mermilliod 1999), making these

effects stronger and consequently enhancing N yields (Meynet & Maeder 2002a,b; Meynet, Ekström & Maeder 2006). We find similar results here comparing various theoretical models for massive star nucleosynthesis (see discussion in § 2.1).

In sufficiently massive AGB stars, the base of the convective envelope is hot enough to activate proton capture reactions, allowing the CNO cycle to convert C and O isotopes into ^{14}N : a process known as hot bottom burning (HBB). AGB stars are also known to experience thermal pulsations, and often these pulsations are accompanied by a penetration of the convective envelope into the CO-rich core, which incorporates some of this material into the envelope itself: a process known as third dredge-up (TDU). When both processes are active, TDU adds new seed nuclei for HBB to turn into ^{14}N , substantially increasing N yields. We demonstrate in §§ 2.2 and 2.3 that various published theoretical models predict significantly discrepant N yields for high mass AGB stars as a consequence of differences in TDU and HBB. The differences in these processes are in turn a consequence of the uncertain microphysical assumptions built into stellar evolution models (e.g. mass loss, opacity, convection and con-

vective boundaries, nuclear reaction networks). In § 4.2, we test the extent to which each of these “off-the-shelf” yield models are able to reproduce the $[\text{N/O}]-[\text{O/H}]$ relation in GCE models.

With a sample of 6,507 galaxies from the MaNGA IFU survey (Bundy et al. 2015), Schaefer et al. (2020) demonstrate that the intrinsic scatter in the $[\text{N/O}]-[\text{O/H}]$ relation at fixed galaxy mass is correlated with variations in the local star formation efficiency (SFE). In regions of slower star formation, $[\text{N/O}]$ tends to be slightly higher at fixed $[\text{O/H}]$ (see their Fig. 4), which is expected from simple GCE models. In classical “closed-box models” (e.g. Mollá et al. 2006), more AGB stars enrich the ISM with N by the time a given $[\text{O/H}]$ is reached, whereas in “open-box models” with inflows and outflows like the ones we present here, dilution by primordial gas accretion drives $[\text{O/H}]$ down at fixed $[\text{N/O}]$. However, Schaefer et al. (2020) did not rule out stellar migration as an additional source of scatter in the gas-phase $[\text{N/O}]-[\text{O/H}]$ relation. In principle, there could be a deficit or surplus of N-producing AGB stars in a given Galactic region at any time simply because the orbits are evolving, driving additional scatter in the correlation. The Johnson et al. (2021) GCE model is an ideal tool with which to test this hypothesis; the novel difference between theirs and previous models with similar motivations is that it allows stellar populations to enrich rings at different radii as they migrate. Originally developed to study the abundances of O and iron (Fe), this aspect of Galactic evolution turned out to have an important impact on the delayed type Ia supernova (SN Ia) enrichment of Fe, causing stochastic fluctuations in the enrichment rates with time at fixed radius. Here we use the same methodology to test for similar effects in the delayed AGB star production of N, in turn assessing whether migration or variability in the SFE dominate scatter in the $[\text{N/O}]-[\text{O/H}]$ relation.

With stellar abundance data, we can test the N abundances predicted by our model against observables unavailable for the gas phase, such as age and $[\text{O/Fe}]$. Using data from the Apache Point Observatory Galaxy Evolution Experiment (APOGEE; Majewski et al.

2017) with asteroseismic mass measurements, Vincenzo et al. (2021) demonstrate that when stellar N abundances are corrected for internal mixing processes, the correlations with stellar age and other elemental abundances are affected. Whether or not our GCE model is able to reproduce their data constitutes a valuable test of our understanding of N nucleosynthesis and the history of N enrichment in the Milky Way. Vincenzo et al. (2021) find good agreement between the APOGEE abundances and the Dopita et al. (2016) data, which we find to be a good representation of external galaxies as well; we therefore take the Dopita et al. (2016) trend (the purple points in Fig. 1) as our observational benchmark.

In § 2, we discuss our adopted yields of N from its dominant nucleosynthetic sources. We discuss the details of our multi-zone chemical evolution model in § 3. We describe the evolution of a fiducial model in § 4.1. In § 4.2, we quantify the $[\text{N/O}]-[\text{O/H}]$ relation predicted by our model with various “off-the-shelf” AGB star yield models taken from the literature. We investigate the relative importance of the delay-time distribution and the metallicity-dependence of AGB star yields in § 4.3. We compare our model predictions to stellar N abundances corrected for internal mixing processes in § 4.4. We assess the sources of intrinsic scatter in the $[\text{N/O}]-[\text{O/H}]$ relation in § 4.5 and summarize our conclusions in § 5.

2 NUCLEOSYNTHESIS

Here we make use of the chemical evolution model for the Milky Way presented by Johnson et al. (2021), which runs using the publicly available *Versatile Integrator for Chemical Evolution* (VICE, see Appendix A; Johnson & Weinberg 2020; Griffith et al. 2021; Johnson et al. 2021), an open-source PYTHON package designed for GCE modeling. Johnson et al. (2021) focus their discussion of the model predictions on O and Fe, and we retain their yields of these elements here. The supernova (SN) yields are defined as the net mass of some element X produced over all explosion events in units of the progenitor cluster’s mass.

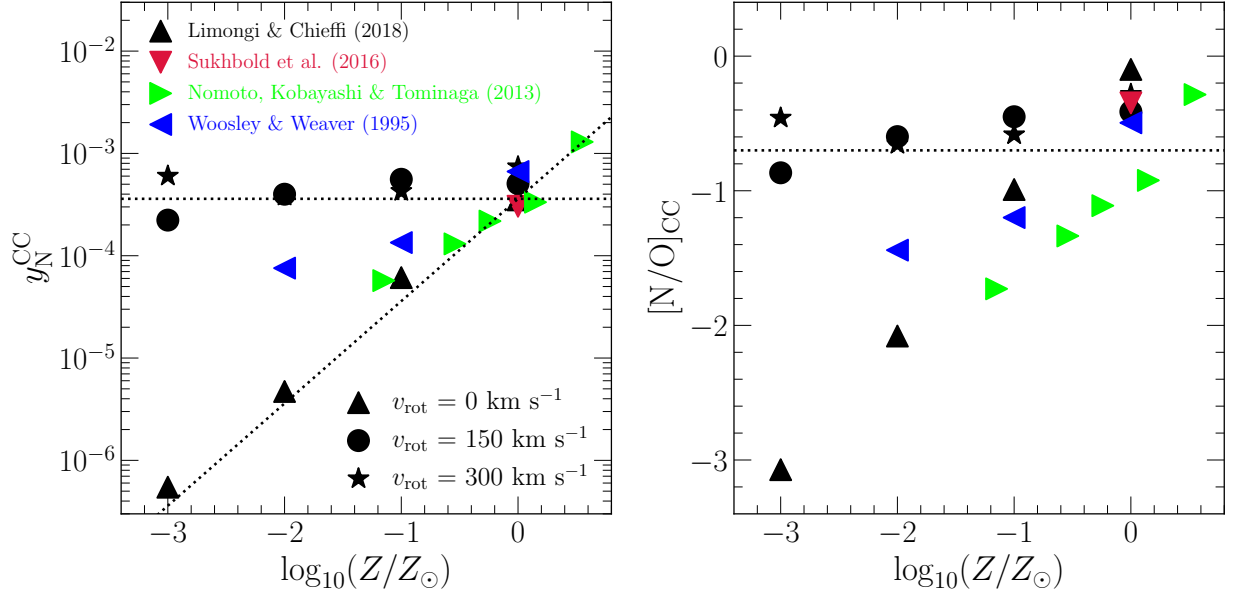


Figure 2. **Left:** IMF-averaged CCSN yields of N calculated using VICE’s `vice.yields.ccsne.fractional` function with the tables published by Woosley & Weaver (1995, blue), Nomoto et al. (2013, green), Sukhbold et al. (2016, red), and Limongi & Chieffi (2018, black). All studies report yields for non-rotating progenitors, shown by the triangles; for visual clarity, the triangles point in a different direction for each study according to the legend. Limongi & Chieffi (2018) report additional yields for progenitors with rotational velocities of 150 (circles) and 300 km/s (stars). The horizontal dashed line marks $y_N^{\text{CC}} = 3.6 \times 10^{-4}$, the value of our fiducial CCSN yield of N in our GCE models. We also use the form shown by the slanted line (equation 8) in § 4.2.2 in combination with some of our AGB star yield models discussed in § 2.2. **Right:** The [N/O] ratio predicted by each of the explosion models in the left-hand panel, under the same colour-coding and marker scheme. We mark the position of [N/O] = -0.7 with a black dotted line, the value roughly suggested by the observations of low-metallicity systems highlighted in Fig. 1.

For example, with a yield of $y_X = 0.001$, a hypothetical $1000 M_{\odot}$ star cluster would produce $1 M_{\odot}$ of the element X instantaneously in the case of CCSNe or over the delay-time distribution (DTD) in the case of SNe Ia. These yields are net yields in that they do not include the metal mass ejected to the interstellar medium (ISM) that was initially present within a star; in the previous example, the $1 M_{\odot}$ yield is only the newly produced metal mass. We adopt the following values from Johnson et al. (2021):

- $y_{\text{O}}^{\text{CC}} = 0.015$
- $y_{\text{Fe}}^{\text{CC}} = 0.0012$
- $y_{\text{O}}^{\text{Ia}} = 0$
- $y_{\text{Fe}}^{\text{Ia}} = 0.00214$,

where the subscripts and superscripts differentiate between the element and the SN type. These choices are based on a mix of theoretical and empirical considerations. For a Kroupa (2001) IMF, the solar metallicity CCSN yields of Limongi & Chieffi (2006) and Griffith

et al. (2021, based on the Sukhbold et al. 2016 models with forced explosion) predict $y_{\text{O}}^{\text{CC}} = 0.017$ and 0.018 , respectively, if all stars from $8 - 120 M_{\odot}$ explode. The value of $y_{\text{O}}^{\text{CC}} = 0.015$ allows for a modest amount of black hole formation but implicitly assumes that most massive stars explode.¹ In chemical evolution models, this choice of yield also leads to good agreement with the observationally inferred deuterium-to-hydrogen ratio of the local interstellar medium (ISM) (Linsky et al. 2006), while substantially lower y_{O}^{CC} leads to disagreement (Weinberg 2017).

Our adopted values of y_{O}^{CC} and $y_{\text{Fe}}^{\text{CC}}$ give $[\text{O}/\text{Fe}] \approx 0.43$ for stars with pure CCSN enrichment, in good agreement with the “high- α ” plateau of disk stars found by Ramírez et al. (2013); matching the APOGEE plateau at $[\text{O}/\text{Fe}] \approx 0.35$ (see, e.g., fig. 6 of Hasselquist et al. 2021) would instead require a slightly higher $y_{\text{Fe}}^{\text{CC}} = 0.0014$.

¹ If all stars from $8 - 40 M_{\odot}$ explode and all more massive stars collapse, then the Sukhbold et al. (2016) models with forced explosions yield $y_{\text{O}}^{\text{CC}} = 0.013$ (Griffith et al. 2021).

SN Ia models predict minimal O yields, justifying $y_{\text{O}}^{\text{Ia}} = 0$. The choice of $y_{\text{Fe}}^{\text{Ia}} = 0.00214$ then leads to good agreement with the observed $[\text{O}/\text{Fe}]$ values of low- α thin-disc stars given the star formation assumptions used by Johnson et al. (2021) (for analytic discussion, see § 3.1 of Weinberg, Andrews & Freudenburg 2017). For an Fe yield of $0.77 M_{\odot}$ from a single SN Ia event (Iwamoto et al. 1999), this $y_{\text{Fe}}^{\text{Ia}}$ corresponds to a time-integrated SN Ia rate of $R_{\text{Ia}} = 2.7 \times 10^{-3} M_{\odot}^{-1}$ (i.e., 2.7 SNe Ia per $1000 M_{\odot}$ of star formation), which is moderately higher than the value of $2.2 \times 10^{-3} M_{\odot}^{-1}$ inferred by Maoz & Mannucci (2012) for a Kroupa (2001) IMF. Our choice of yields is internally consistent and reproduces many Milky Way observations (Johnson et al. 2021), but many of the GCE model predictions would be minimally affected if we lowered y_{O}^{CC} , $y_{\text{Fe}}^{\text{CC}}$, and $y_{\text{Fe}}^{\text{Ia}}$ by a common factor and reduced the efficiency of outflows. We return to this point in the context of N yields in § 4.2.1.

We assume that N is not produced in significant amounts by SNe Ia (Johnson 2019), setting $y_{\text{N}}^{\text{Ia}} = 0$. The remainder of this section discusses the CCSN and AGB star yields of N.

A significant portion of N yields arise as a consequence of the CNO cycle.² As the dominant source of pressure and energy generation in non-zero metallicity main sequence stars with initial masses of $\gtrsim 1.3 M_{\odot}$, this cyclic nuclear reaction catalyses the conversion of H into He that would otherwise be accomplished by the proton-proton chain (von Weizsäcker 1937, 1938; Bethe 1939a,b; Adelberger et al. 2011; Suliga, Shalgar & Fuller 2021). Its slowest component by far is the $^{14}\text{N}(\text{p}, \gamma)^{15}\text{O}$ reaction (e.g. LUNA Collaboration et al. 2006). Consequently, the first order effect of the CNO cycle is to convert most of the C isotopes in stellar cores into ^{14}N . As we will discuss in this section, this plays an important role in shaping N yields from stars of all masses.

² $^{12}\text{C}(\text{p}, \gamma)^{13}\text{N}(\beta^+ \nu_e)^{13}\text{C}(\text{p}, \gamma)^{14}\text{N}(\text{p}, \gamma)^{15}\text{O}(\beta^+ \nu_e)^{15}\text{N}(\text{p}, \alpha)^{12}\text{C}$

2.1 Core Collapse Supernovae and Massive Star Winds

In VICE, CCSN nucleosynthetic products are approximated to be produced instantaneously following an episode of star formation; this is a good approximation because the lives of massive stars are short compared to the relevant timescales for GCE. The yield is simply the constant of proportionality between the CCSN production rate and the star formation rate (SFR):

$$\dot{M}_{\text{X}}^{\text{CC}} = y_{\text{X}}^{\text{CC}} \dot{M}_{\star} \quad (1)$$

More generally, y_{X}^{CC} quantifies *all* of the nucleosynthetic material approximated to be produced instantaneously following a single stellar population’s formation, including newly synthesized material expelled in a massive star wind before the star explodes or collapses to a black hole.

We compute theoretically predicted values of y_{N}^{CC} using VICE’s `vice.yields.ccsne.fractional` function assuming a Kroupa (2001) IMF; details on how VICE handles these calculations can be found in § 4 of Griffith et al. (2021) and in the VICE science documentation³. In the left panel of Fig. 2, we plot the results as a function of progenitor metallicity as predicted by the Woosley & Weaver (1995), Nomoto, Kobayashi & Tominaga (2013), Sukhbold et al. (2016), and Limongi & Chieffi (2018) tables. There is generally good agreement between the various non-rotating models, but only Limongi & Chieffi (2018) report yields for progenitors with non-zero rotational velocities; these yields are substantially larger than their non-rotating counterparts, especially at low metallicity. With few seed nuclei for the CNO cycle at low Z , production of ^{14}N is difficult. Rotation-induced mixing, a highly uncertain process (Zahn 1992; Maeder & Zahn 1998; Lagarde et al. 2012), could transport newly produced C and O into the hydrogen burning shell of the CCSN progenitor, facilitating ^{14}N production (Frischknecht et al. 2016; see also discussion in § 4.2 of Andrews et al. 2017). Consequently, N

³ https://vice-astro.readthedocs.io/en/latest/science_documentation/yields

yields at low metallicity are quite sensitive to model-dependent assumptions regarding stellar rotation and internal mixing processes (Heger & Woosley 2010).

We compute the $[N/O]$ ratio of CCSN ejecta from the values of y_N^{CC} and y_O^{CC} predicted by a given yield table according to

$$[N/O]_{cc} = \log_{10} \left(\frac{y_N^{CC}}{y_O^{CC}} \right) - \log_{10} \left(\frac{Z_{N,\odot}}{Z_{O,\odot}} \right), \quad (2)$$

where $Z_{X,\odot}$ is the abundance by mass of some element X in the sun, for which we take $Z_{N,\odot} = 6.91 \times 10^{-4}$ and $Z_{O,\odot} = 5.72 \times 10^{-3}$ based on the photospheric measurements of Asplund et al. (2009). For each value of y_N^{CC} in the left panel of Fig. 2, we compute the corresponding values of y_O^{CC} and illustrate the resultant $[N/O]_{cc}$ ratios in the right panel. These yield ratios follow similar trends with progenitor metallicity and rotation as y_N^{CC} itself, a consequence of the fact that these studies predict relatively metallicity- and rotation-independent O yields. At low metallicity, CCSN yields of N dominate over the AGB star yields (see discussion in § 2.2), and Fig. 1 suggests a plateau in $[N/O]$ at low metallicity at $[N/O]_{cc} \approx -0.7$. Taking this value in combination with our adopted O yield of $y_O^{CC} = 0.015$, equation (2) suggests that $y_N^{CC} = 3.6 \times 10^{-4}$. We highlight both $[N/O]_{cc} = -0.7$ and $y_N^{CC} = 3.6 \times 10^{-4}$ with horizontal black dashed lines in Fig. 2, finding good agreement with the rotating progenitor models of Limongi & Chieffi (2018) in both panels. This indicates that rotating massive stars play an important role in establishing the N abundances at low metallicity, in agreement with previous works (Chiappini et al. 2003, 2005, 2006; Kobayashi et al. 2011; Prantzos et al. 2018; Grisoni et al. 2021). We therefore take $y_N^{CC} = 3.6 \times 10^{-4}$ as our fiducial CCSN yield of N; both the normalization and metallicity-independence of this choice are supported by the Limongi & Chieffi (2018) models.

The Sukhbold et al. (2016) tables, available only at solar metallicity, agree nearly perfectly with our empirical value of $y_N^{CC} = 3.6 \times 10^{-4}$, but they predict a higher value of $[N/O]_{cc}$ by ~ 0.2 dex. This is a consequence of the failed supernovae incorporated into their model and the lowered values of y_O^{CC} that result (see discussion

in § 4.2). While N emerges in substantial amounts in winds, much of the O produced by massive stars is ejected during the explosion, making the O yield more sensitive to the black hole landscape (Griffith et al. 2021). Most of the SN models plotted in Fig. 2 estimate slightly higher $[N/O]_{cc}$ at $\log_{10}(Z/Z_{\odot}) = 0$ relative to our empirical value of $[N/O]_{cc} = -0.7$, but they still fall short of solar $[N/O]$. This implies the need for an additional enrichment channel, which is expected because it is well understood that N is also produced in considerable amounts by AGB stars (e.g. Johnson 2019).

2.2 Asymptotic Giant Branch Stars

Similar to SNe, our AGB star yields are parametrized as fractional net yields. For a yield y_X^{AGB} , the mass yield is given by $M_{\star} y_X^{AGB}$. Enrichment proceeds as it does in Johnson et al. (2021): AGB stars place their nucleosynthetic products in the $\delta R_{gal} = 100$ pc ring that they are in at a given time, allowing stars to enrich distributions of radii as they migrate. VICE implements an algorithm that computes the mass in dying stars from each stellar population, and the zero age main sequence (ZAMS) mass required to compute the fractional yield comes from a mass-lifetime relationship. For the latter, we adopt the metallicity independent parabola in $\log \tau - \log m$ space from Larson (1974) with updated coefficients from Kobayashi (2004) and David, Forman & Jones (1990) (see discussion of the mass-lifetime relationship in VICE in Appendix A).

We make use of four previously published tables of AGB star N yields computed from stellar evolution models, each of which are sampled on a grid of progenitor masses and metallicities. To approximate the net yield y_X^{AGB} as a smooth function of M_{\star} and Z_{\star} , VICE interpolates bi-linearly - once in mass M and once in metallicity Z - and linearly extrapolates above or below the grid in either quantity as necessary. By comparing the predicted abundances of the Johnson et al. (2021) Milky Way model to the latest observational data, we can constrain how accurately these “off-the-shelf”

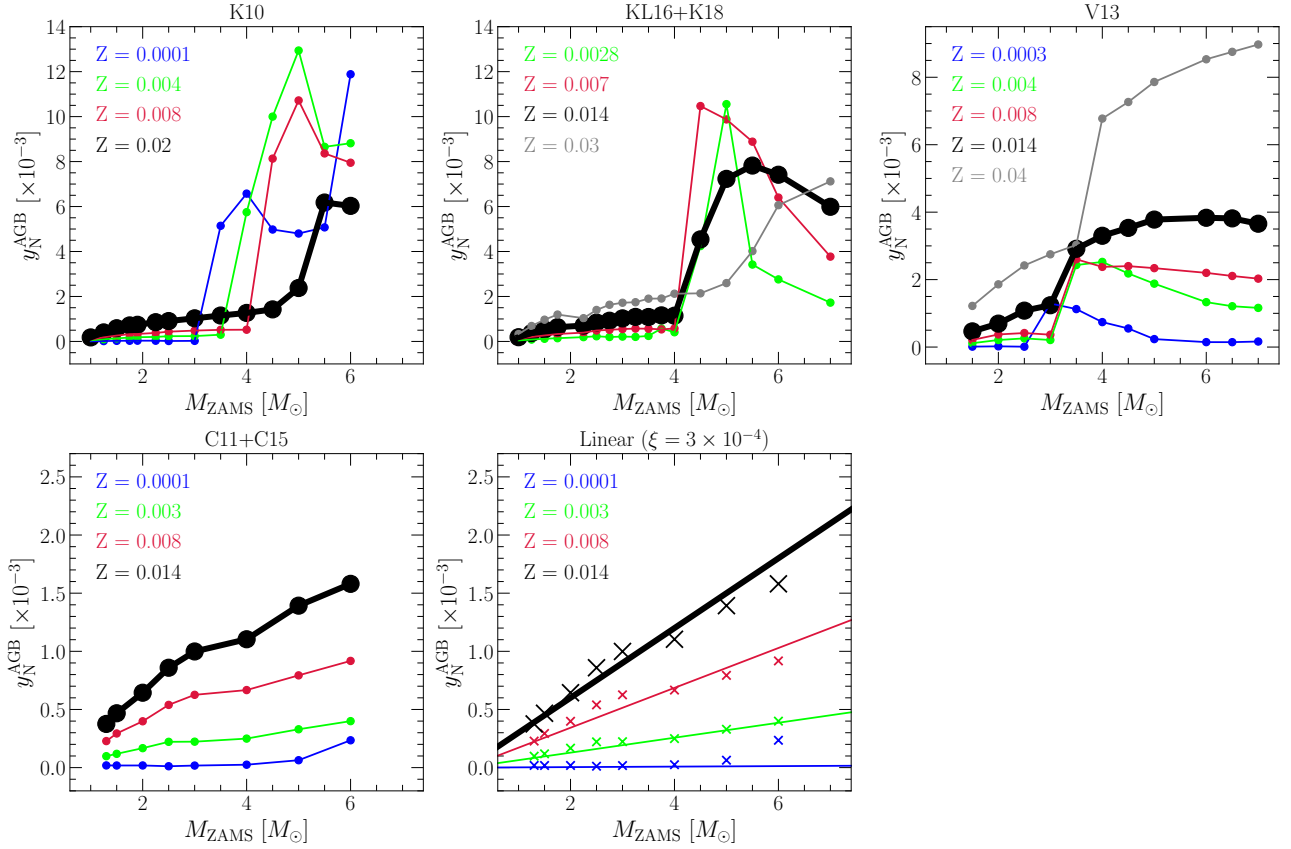


Figure 3. The fractional yields of N from AGB stars y_N^{AGB} as a function of progenitor ZAMS mass and birth metallicity Z as reported by Karakas (2010, upper left), Karakas & Lugaro (2016) and Karakas et al. (2018, upper middle), Ventura et al. (2013, 2014, 2018, 2020, upper right), and Cristallo et al. (2011, 2015, lower left). For Ventura et al. (2013, 2014, 2018, 2020) and Cristallo et al. (2011, 2015), we show the yields only for a selection of metallicities available from their provided tables. We highlight yields at solar metallicity ($Z = 0.02$ for Karakas 2010, $Z = 0.014$ otherwise) with bold black lines. In the lower right panel, we show our linear model (coloured lines, see equation 3) in comparison to the Cristallo et al. (2011, 2015, coloured X's) predictions. We caution that the y-axis ranges are not the same between panels in this figure.

yield models characterize N production. These models taken from the literature are as follows.

1. Karakas (2010, hereafter K10)⁴ published yields for $Z = 0.0001, 0.004, 0.008$, and 0.02 progenitors. We plot these yields in the upper left panel of Fig. 3.
2. Karakas & Lugaro (2016) and Karakas et al. (2018) published yields for $Z = 0.0028, 0.007, 0.014$, and 0.03 progenitors; we hereafter refer to these yields as the KL16+K18 model. We illustrate these yields in the upper middle panel of Fig. 3.
3. We combine the yields for $Z = 0.0003$ and 0.008 progenitors

⁴ We clarify that our abbreviations of each of these papers refer specifically to their yields of N as we adopt them in our model. We cite the full names of these papers when referring to their more general results.

from Ventura et al. (2013) with those at $Z = 0.004$ from Ventura et al. (2014), at $Z = 0.014$ from Ventura et al. (2018), and at $Z = 0.04$ from Ventura et al. (2020) into a single table of yields. In this set, we also include a set of un-published yields at $Z = 0.001$ and 0.002 computed from similar models (provided by P. Ventura, private communication). We hereafter refer to this yield set as the V13 model, and we plot a subsample of these yields in the upper right panel of Fig. 3.

4. The default set of AGB star yields in VICE is taken from Cristallo et al. (2011, 2015), who published yields for $Z = 0.0001, 0.0003, 0.001, 0.002, 0.003, 0.006, 0.008, 0.01, 0.014$, and 0.02 progenitors. We hereafter refer to these yields as the C11+C15 model, and we illustrate a subsample of them in the lower left panel of Fig. 3.

VICE also allows users to construct their own functions of progenitor mass and metallicity to describe the AGB star yield. Motivated by the roughly linear nature of the C11+C15 yields and their general success once renormalized by a constant factor (see discussion in § 4.2), we construct a model in which the yield is linearly proportional to both progenitor ZAMS mass and metallicity:

$$y_N^{\text{AGB}} = \xi \left(\frac{M_\star}{M_\odot} \right) \left(\frac{Z_\star}{Z_\odot} \right). \quad (3)$$

We illustrate this model in the lower middle panel of Fig. 3 for $\xi = 3 \times 10^{-4}$ in comparison to the C11+C15 yields shown by the coloured X's. Although we find good agreement between the C11+C15 yields and our linear model with a normalization of $\xi = 3 \times 10^{-4}$, for our fiducial AGB star yield of N we take a slope of $\xi = 9 \times 10^{-4}$. We discuss the absolute scaling of our nucleosynthetic yields in § 4.2 below.

As is clear from Fig. 3, the N yields reported by these studies show substantial differences. Unfortunately, ascertaining the origin of these differences is difficult because each model employs its own assumptions for important evolutionary parameters such as opacity, mass loss, nuclear reaction networks, and convection and convective boundaries within stars, all of which have a significant impact on stellar evolution and thus the predicted yields (see discussion in, e.g., § 5 of Karakas & Lugaro 2016). However, the differences can be qualitatively understood by considering two important phenomena known to occur within AGB stars: TDU⁵ and HBB. The variations in how TDU and HBB proceed between different stellar evolution models arise as consequences of the different input physics.

When an AGB star experiences a thermal pulsation, this is usually accompanied by a TDU event whereby the convective envelope penetrates into the hydrogen-depleted core, mixing some of this material with other material exposed to partial He-shell burning. The $^{13}\text{C}(\alpha,$

⁵ The time adverbial “third” in TDU refers only to the fact that these dredge-up episodes are occurring while the star is on the asymptotic giant branch. Because they are associated with the thermal pulsations of AGB stars, there are many episodes of third dredge-up.

$n)^{16}\text{O}$ reaction, which is the main source of free neutrons in low-mass AGB stars (Gallino et al. 1998), can occur at substantial rates when this core material is mixed with the He-rich shell. This process does not directly affect N abundances in the shell because the core is mostly composed of C and O at this evolutionary phase, but ^{14}N plays an important role in shaping an AGB star's overall s-process yield by acting as an efficient catalyst of neutron decay via the $^{14}\text{N}(n, p)^{14}\text{C}(\beta^+ \nu_e)^{14}\text{N}$ reaction, the first step of which is a resonant neutron capture (Cristallo et al. 2011).

HBB refers to proton capture reactions at the base of the convective envelope, activating the CNO cycle and producing large amounts of ^{14}N at the expense of C and O isotopes. HBB requires a higher mass AGB star progenitor ($M_{\text{ZAMS}} = 4 - 5 M_\odot$ at Z_\odot according to Karakas 2010) than TDU ($M_{\text{ZAMS}} = 2 - 2.5 M_\odot$ at Z_\odot according to Karakas 2010), but the minimum mass for both decreases at lower metallicities.

The most efficient N production occurs when both TDU and HBB are active within an AGB star, because each replenishment of C and O isotopes by TDU adds new seed nuclei for the CNO cycle with HBB. This is the reason for the substantial increase in yields at $\sim 4 M_\odot$ in the K10 and KL16+K18 models; in both yield sets, every star that experiences HBB also experiences TDU (see, e.g., Table 1 of Karakas 2010). Their high mass AGB star yields are higher at low Z because both HBB and TDU are more efficient (see discussion in Ventura et al. 2013): when the metallicity is low, each TDU episode is deeper due to the lower opacity, and the base of the convective envelope is hotter, increasing the rate of CNO cycle reactions in HBB. This interaction between TDU and HBB is also the reason for the increase in the V13 yields near $\sim 3 M_\odot$, but unlike the K10 and KL16+K18 models, their stars experience both processes only in this narrow range of mass.

Of all of these yields taken from the literature, the C11+C15 sample shows the smoothest dependence on progenitor mass and metallicity. Below $\sim 3 M_\odot$, their agreement with the KL16+K18 yields is

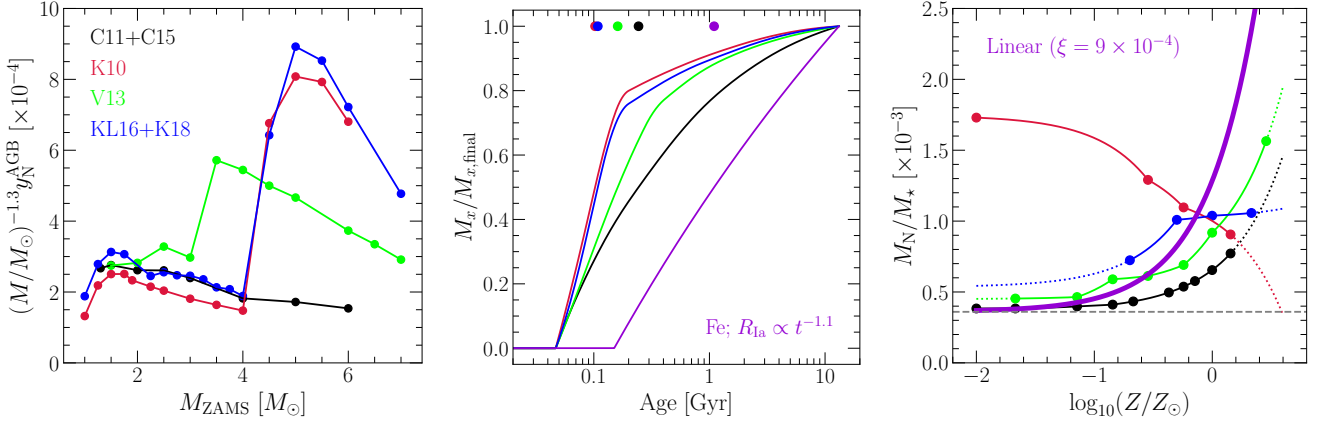


Figure 4. **Left:** The IMF-weighted mass yield of N from AGB stars as a function of progenitor ZAMS mass at solar metallicity (i.e. the contribution per linear interval dM_{ZAMS} ; $Z_{\odot} = 0.014$). **Middle:** The net mass of N produced by AGB stars from a single stellar population for each of our yield models at solar metallicity. The purple line denotes the same for Fe assuming our $t^{-1.1}$ DTD as in the Johnson et al. (2021) chemical evolution model. All values are normalized to the total mass produced at an age of 13.2 Gyr. Points at the top of the panel denote the ages at which 50% of the total mass yield has been produced. **Right:** The total amount of N produced by a 13.2 Gyr old stellar population as a function of metallicity for each of our yield models normalized by the stellar population’s initial mass. Points mark metallicities at which the published tables report yields, and the lines are dotted at metallicities that are above (below) the maximum (minimum) metallicity reported by a given study (i.e. where extrapolation is necessary). In this panel only, we include the metallicity-independent contribution $y_N^{\text{CC}} = 3.6 \times 10^{-4}$ from CCSNe (gray dashed line). The bold purple curve represents our inference of the *total* N yield (CCSN + AGB) required to reproduce the observational constraints discussed in § 4 given our adopted O and Fe yields (§ 2).

good, but this model has much lower N yields for higher mass AGB stars. Pinpointing a single reason for this difference is difficult, even when considering the differences between HBB and TDU. Relative to the KL16+K18 yields (see discussion in § 5 of Karakas & Lugaro 2016), the C11+C15 stars have more mass loss, fewer thermal pulses overall, and weaker HBB due to a lower temperature at the base of the convective envelope, each of which acts to lower the yield of ^{14}N .

Although the K10 and KL16+K18 yield models both show a substantial increase in N yields above $\sim 4 M_{\odot}$, there are some noteworthy differences between the two. In the newer version, the yields at solar metallicity are somewhat higher, and the yields at sub-solar metallicities decreased slightly, particularly for the highest mass AGB stars. These differences can be understood by slight variations in the input physics (A. Karakas, private communication). A portion of the increase in the yields at solar metallicity can be attributed to the

assumption of $Z_{\odot} = 0.02$ versus $Z_{\odot} = 0.014^6$ and the impact this has on both HBB and TDU, but it does not account for the entire difference. As a consequence of updates to opacity tables and the adopted solar composition, the KL16+K18 models at solar metallicity are slightly hotter and more compact, giving them hotter HBB and deeper TDU. With more thermal pulses overall and therefore a longer AGB lifetime, these stars have more time to convert ^{12}C into ^{14}N . KL16+K18 also use low-temperature opacity tables based on Marigo (2002) that more closely follow the surface composition of the star. These opacities are high, making the more massive AGB stars larger and increasing the mass-loss rate relative to K10, truncating their N yields. The $Z = 0.0028$ model uses the Bloeker (1995) mass-loss prescription rather than that of Vassiliadis & Wood (1993), which was used for the K10 yields as well as the yields at other metallicities in the KL16+K18 model. This choice results in fewer thermal pulses

⁶ Changes in the accepted value of the metallicity of the sun trace back to the canonical value of $\sim 2\%$ derived by, e.g., Anders & Grevesse (1989) and Grevesse & Sauval (1998), later being revised to $\sim 1.4\%$ by, e.g., Lodders (2003) and Asplund, Grevesse & Sauval (2005). See Table 4 of Asplund et al. (2009) for a compilation of measured values.

and a shorter AGB lifetime, giving them less time to process C and O nuclei into ^{14}N .

In the interest of consistency, when we adopt a particular AGB star yield model for N, we also adopt the corresponding table within VICE for O and Fe when possible.⁷ These O and Fe yields, however, are negligible compared to their SN yields. As a reference for the reader, in Table 1 we provide a summary of our AGB star yield models, where they are defined or published, and in what sections we present GCE model predictions with each yield. Although we focus our investigation on AGB yields from $\lesssim 7 M_{\odot}$ stars, slightly more massive stars (up to $\sim 12 M_{\odot}$) sit near the critical mass boundaries between different types of massive white dwarfs and electron capture SN progenitors. Doherty et al. (2017) investigated theoretically predicted yields of these stars and found significant production of CNO isotopes. There is also the intriguing possibility of the CNO yields from the earliest, most metal-poor AGB stars (e.g. the $Z = 10^{-5}$ models of Gil-Pons et al. 2013, 2021) and the insight this may afford into N production at low Z and the most metal-poor stars in the Galaxy. While experiments with such yields in our GCE models would be interesting, this is beyond the scope of the current paper since our AGB yield models already span a wide range of assumptions regarding stellar evolution.

2.3 IMF-Averaged AGB Star Yields: Metallicity and Time Dependence

To more directly compare these AGB star yields predicted from stellar evolution models, we plot their IMF-weighted yields at solar metallicity in the left hand panel of Fig. 4. We assume $Z_{\odot} = 0.014$ based on Asplund et al. (2009) and Asplund, Amarsi & Grevesse (2021); since the K10 model reports yields at $Z = 0.02$ rather than $Z = 0.014$,

⁷ In the case of Ventura et al. (2013, 2014, 2018, 2020), AGB star yields of Fe are not available, and our linear model is only appropriate for N. In these cases, we assume the VICE default of the Cristallo et al. (2011, 2015) yields for both O and Fe.

we simply interpolate linearly to $Z = 0.014$ in the same manner that VICE does in our GCE models. As mentioned in § 2.2, the AGB star yield $y_{\text{N}}^{\text{AGB}}$ as we have parametrized it is in units of the progenitor star’s ZAMS mass, and consequently the mass yield of N is given by $M_{\star} y_{\text{N}}^{\text{AGB}}$. With an additional weight of $M_{\star}^{-2.3}$ from the IMF in this mass range (e.g. Kroupa 2001), we therefore multiply the values of $y_{\text{N}}^{\text{AGB}}$ by $(M_{\star}/M_{\odot})^{-1.3}$ to quantify a star’s relative contribution to the total N yield taking into account the intrinsic mass distribution.⁸ With the additional weight of $M_{\star}^{-1.3}$, the C11+C15 yields are relatively mass-independent. For the other yield models, higher mass AGB stars dominate the overall yield due to the effects of TDU and HBB discussed in § 2.2.

Using VICE’s `vice.single_stellar_population` function, in the middle panel of Fig. 4 we plot the total N yield as a function of age from a single stellar population. For the sake of this calculation, we set all CCSN yields of N to zero in order to highlight the AGB star contribution. We show the results of this procedure for solar metallicity only (again $Z = 0.02$ for K10, $Z = 0.014$ otherwise), and we normalize all values to the total mass produced at $T = 13.2$ Gyr (the total amount of time our GCE model is integrated over; see discussion in § 3). Under the C11+C15 yields, it takes ~ 250 Myr for a single stellar population to produce $\sim 50\%$ of its N from AGB stars, as noted by the coloured points at the top of the panel. This is in good agreement with Maiolino & Mannucci (2019), who find that similar parameter choices predict 80% of the N yield to be ejected within ~ 1 Gyr (see their Fig 1). The characteristic timescales for N production are even shorter in the other yield models because of their more pronounced contributions from massive stars with short lifetimes (e.g. Larson 1974; Maeder & Meynet 1989; Padovani & Matteucci 1993). For comparison, we plot the enrichment of Fe by our $t^{-1.1}$ power-law DTD, also with the CCSN yield set to zero to highlight the delayed component. The characteristic delay time for

⁸ This weight gives a contribution per linear interval of M_{ZAMS} , so one can use area under the curve to assess relative contributions.

Table 1. A summary of our AGB star yield models, where they are published/defined, and in which sections we present tests of each model. We take a fiducial value of $y_N^{\text{CC}} = 3.6 \times 10^{-4}$ for our IMF-averaged massive star yield (see discussion in § 2.1), and we explore alternate parametrizations of this value in § 4.2.2 only (right panel of Fig. 6).

Model	Reference/Definition	Results
Linear	equation (3) (fiducial yield)	4.1, 4.2.1, 4.3, 4.4, 4.5
C11+C15	Cristallo et al. (2011, 2015)	4.2, 4.2.1
V13	Ventura et al. (2013, 2014, 2018, 2020)	4.2, 4.2.1
K10	Karakas (2010)	4.2, 4.2.2
KL16+K18	Karakas & Lugaro (2016); Karakas et al. (2018)	4.2, 4.2.2

Fe production is considerably longer than that of N - up to an order of magnitude depending on which yield model is adopted. As noted in Johnson et al. (2021), a characteristic delay time of ~ 1 Gyr is exactly as expected for a $\sim t^{-1}$ DTD because half of the SN Ia events occur between 100 Myr and 1 Gyr and the other half between 1 Gyr and 10 Gyr.

A characteristic delay time of only ~ 250 Myr may seem surprising given the relatively mass-independent nature of the IMF-weighted C11+C15 yields. This arises out of the steep nature of the stellar mass-lifetime relation (e.g. Larson 1974; Maeder & Meynet 1989; Padovani & Matteucci 1993). For example, 2 and 3 M_\odot stars live only ~ 1.2 Gyr and ~ 400 Myr, respectively, and over the course of 13.2 Gyr, only stars of masses $\gtrsim 0.9 M_\odot$ will have enough time to finish their hydrogen burning. Consequently, most of the mass range of stars that will evolve through an AGB phase will do so within the first few hundred Myr after their formation, and with mass-independent IMF-weighted yields, this accounts for most of the N. We clarify that the delay times computed here apply *only* to N and not necessarily to other elements produced by AGB stars. As we have illustrated here, the effective DTD of AGB star enrichment is dictated by the combination of the stellar mass-lifetime relation and the mass dependence of the yield, which should in principle differ from element to element. Other elements produced by slow neutron capture often have the highest yields from lower mass AGB stars. For example, Cristallo et al. (2011, 2015) report Sr yields that are dominated by $M_{\text{ZAMS}} = 2 - 3 M_\odot$ progenitors (see fig. 5 of Johnson & Weinberg 2020), giving it a characteristic delay time of ~ 500 Myr.

The characteristic delay-times will be as long as a few Gyr if and only if the yields are dominated by $\lesssim 1.5 M_\odot$ stars.

In the right panel of Fig. 4, we plot the total amount of N produced by a 13.2 Gyr old single stellar population as a function of its initial metallicity according to all of our AGB star yield tables, including the linear model (see equation 3 and discussion in § 2.2). For this calculation, we include the metallicity independent CCSN yield ($y_N^{\text{CC}} = 3.6 \times 10^{-4}$; see discussion in § 2.1). In general, there is good qualitative agreement between the C11+C15 and the V13 models, the only major difference being the normalization. The predictions with the linear model with $\xi = 3 \times 10^{-4}$ are nearly identical to the C11+C15 model, as one would expect from Fig. 3, but here we show the yields for our fiducial choice of $\xi = 9 \times 10^{-4}$. The value at which these N yields flatten off at low Z is reflective of our adopted value of y_N^{CC} . Up to $\log_{10}(Z/Z_\odot) \approx -0.2$, the KL16+K18 yields predict a similar trend as C11+C15 and V13, also with a difference in normalization, but at solar and super-solar metallicities they predict much more metallicity-independent N yields than others. The K10 yields, on the other hand, do not agree with any of the other models, instead predicting N yields to *decrease* monotonically with increasing Z . These differences between the K10 and KL16+K18 models trace back to differences regarding the opacity and mass loss prescriptions (see discussion in § 2.2). Although the normalization depends on the SN yields of all elements, we demonstrate in § 4.2 that reproducing the [N/O]-[O/H] relation as observed requires AGB N yields which scale roughly linearly with metallicity as in the C11+C15 and V13 models. More specifically, with our adopted O and Fe yields (see discussion at the beginning of § 2), reproducing the observational constraints that we consider requires *total* N yields (CCSN + AGB) with the metallicity dependence shown by the purple curve in the right panel of Fig. 4.

3 THE MULTI-ZONE CHEMICAL EVOLUTION MODEL

We use the fiducial model for the Milky Way published by [Johnson et al. \(2021\)](#), which runs using the VICE GCE code (see Appendix A; [Johnson & Weinberg 2020](#); [Griffith et al. 2021](#)). Multi-zone models allow simultaneous calculations of abundances for multiple Galactic regions, making them a more physically realistic option than classical one-zone models for a system like the Milky Way. Furthermore, they can take into account stellar migration in a framework that is much less computationally expensive than hydrodynamical simulations, making them the ideal experiments for our purposes. We provide a brief summary of the model here, but a full breakdown can be found in § 2 of [Johnson et al. \(2021\)](#).

As in previous models for the Milky Way with similar motivations (e.g. [Matteucci & Francois 1989](#); [Schönrich & Binney 2009](#); [Minchev et al. 2013, 2014, 2017](#); [Sharma et al. 2021](#)), this model parametrizes the Galaxy disc as a series of concentric rings. With a uniform width of $\delta R_{\text{gal}} = 100$ pc, each ring is assigned its own star formation history (SFH), and with assumptions about outflows and the $\Sigma_{\text{gas}} - \dot{\Sigma}_{\star}$ relation (see discussion below), VICE computes the implied amounts of gas and infall at each timestep automatically. Under the caveat that stellar populations can move and place some of their nucleosynthetic products in rings other than the one they were born in, each ring is otherwise described by a conventional one-zone GCE model. Allowing stars to enrich distributions of radii was a novel addition to this type of model, and [Johnson et al. \(2021\)](#) demonstrated that this has a significant impact on enrichment rates from delayed sources such as SNe Ia.

To drive stellar migration, the model makes use of star particles from a hydrodynamical simulation, for which [Johnson et al. \(2021\)](#) select the h277 galaxy from the [Christensen et al. \(2012\)](#) suite evolved with the N-body+SPH code GASOLINE ([Wadsley et al. 2004](#)); we retain this decision here. h277 spans 13.7 Gyr of evolution, but the sample of star particles with reliable birth radii span 13.2 Gyr in age; the model thus places the onset of star formation ~ 500 Myr after

the Big Bang and integrates up to the present day. Previous studies have shown that h277, among other disc galaxies evolved with similar physics, has a realistic rotation curve ([Governato et al. 2012](#); [Christensen et al. 2014a,b](#)), stellar mass ([Munshi et al. 2013](#)), metallicity ([Christensen et al. 2016](#)), dwarf satellite population ([Zolotov et al. 2012](#); [Brooks & Zolotov 2014](#)), HI properties ([Brooks et al. 2017](#)), and stellar age-velocity relation ([Bird et al. 2021](#)). Despite this, there are some interesting differences between h277 and the Milky Way. First and foremost, h277 had only a weak and transient bar and lacks one at the present day, while the Milky Way is known to have a strong, long-lived central bar (e.g. [Bovy et al. 2019](#)). This could indicate that the dynamical history of h277 and its star particles differs significantly from that of the Milky Way. Furthermore, the last major merger in h277 was at a redshift of $z \approx 3$, making it an interesting case study for its quiescent merger history (e.g. [Zolotov et al. 2012](#)), while the Sagittarius dwarf galaxy is believed to have made pericentric passages around the Milky Way at 1 – 2 Gyr intervals ([Law & Majewski 2010](#)). Although these differences between h277 and the Milky Way are well understood, their impact on chemical evolution is not. We are unaware of any studies that investigate the impact of different assumptions regarding the Galaxy’s dynamical history on predicted abundances; this is however an interesting question for future work.

Radial migration of stars proceeds from the h277 star particles in a simple manner; for a stellar population in our model born at a radius R_{birth} and a time T_{birth} , VICE searches for star particles born at $R_{\text{birth}} \pm 250$ pc and $T_{\text{birth}} \pm 250$ Myr. From the star particles that make this cut, it then randomly selects one to act as that stellar population’s *analogue*. The stellar population then assumes the present day midplane distance z and the change in orbital radius ΔR_{gal} of its analogue between its birth and the present day. In the [Johnson et al. \(2021\)](#) fiducial model, stellar populations move to their implied final radii with a $\sqrt{\text{age}}$ dependence according to:

$$R_{\text{gal}}(T) = R_{\text{birth}} + \Delta R_{\text{gal}} \sqrt{\frac{T - T_{\text{birth}}}{13.2 \text{ Gyr} - T_{\text{birth}}}}, \quad (4)$$

where 13.2 Gyr is simply the present day (see discussion above). With displacement proportional to \sqrt{age} , this corresponds to a scenario in which radial migration proceeds as a diffusion process as modeled by Frankel et al. (2018, 2019) and supported by the N-body simulations of Brunetti, Chiappini & Pfenniger (2011). Although Johnson et al. (2021) investigated other assumptions for this time-dependence, in the present paper we only use this parametrization (hereafter referred to as the “diffusion” model) and an idealized one in which stars remain at their birth radius until they instantaneously migrate at the present day (hereafter referred to as the “post-processing” model). If VICE does not find any star particles from h277 in its initial $R_{\text{gal}} \pm 250$ pc and $T \pm 250$ Myr search, it widens it to $R_{\text{gal}} \pm 500$ pc and $T \pm 500$ Myr; if still no candidate analogues are found, VICE maintains the $T \pm 500$ Myr requirement, but assigns the star particle with the smallest difference in birth radius as the analogue. This procedure can be thought of as “injecting” the dynamics of the h277 galaxy into our multi-zone chemical evolution model, and it can in principle be repeated for any hydrodynamical simulation of a disc galaxy. As in Johnson et al. (2021), we neglect radial gas flows (e.g. Lacey & Fall 1985; Bilitewski & Schönrich 2012; Vincenzo & Kobayashi 2020), instead focusing on the impact of stellar migration.

Rather than using a hydrodynamical simulation, some previous studies have implemented stellar migration using dynamical arguments (e.g. Schönrich & Binney 2009; Sharma et al. 2021). An advantage of our approach over this is that these dynamical arguments introduce free parameters which then require fitting to data. It is also unclear to what extent the fit might bias the model into agreement with quantities in the sample not involved in the fit. A disadvantage of our approach is that we are restricted to one realization of our dynamical history; slight variations are not possible. We do not distinguish between “blurring” and “churning”, terms commonly used to refer to changes in the guiding centre radii of stars and their epicyclic motions, respectively. Both are induced by a variety of physical interactions such as molecular cloud scattering (Mihalas

& Binney 1981; Jenkins & Binney 1990; Jenkins 1992), orbital resonances with spiral arms or bars (Sellwood & Binney 2002; Minchev et al. 2011), and satellite perturbations (Bird, Kazantzidis & Weinberg 2012). All of these effects are included in h277 and should therefore be inherited by the stellar populations in our GCE model.

We assume the SFH of the “inside-out” model from Johnson et al. (2021). The time-dependence at a given R_{gal} is described by

$$\dot{\Sigma}_{\star} \propto (1 - e^{-t/\tau_{\text{rise}}})e^{-t/\tau_{\text{sfr}}}, \quad (5)$$

where τ_{rise} approximately controls the amount of time the SFR is rising at early times; we set this parameter equal to 2 Gyr at all radii as in Johnson et al. (2021). Our e-folding timescales of τ_{sfr} are taken from a fit of this functional form to the Σ_{\star} -age relation in bins of R/R_{e} for $10^{10.5} - 10^{11} M_{\odot}$ Sa/Sb Hubble type spiral galaxies reported by Sánchez (2020). The resulting values of τ_{sfr} are long: ~ 15 Gyr at the solar circle ($R_{\text{gal}} = 8$ kpc) and as high as ~ 40 Gyr in the outer disc (see fig. 3 of Johnson et al. 2021). This is a consequence of flat nature of the Σ_{\star} -age relation reported by Sánchez (2020).

Within each $\delta R_{\text{gal}} = 100$ pc ring, the normalization of the SFH is set by the total stellar mass of the Milky Way disc and the present-day stellar surface density gradient, assuming that it is unaffected by migration (see Appendix B of Johnson et al. 2021). For the former, we neglect the contribution from the bulge and adopt the total disc stellar mass of $5.17 \times 10^{10} M_{\odot}$ from Licquia & Newman (2015). For the latter, we adopt a double exponential form describing the thin- and thick-disc components. We take the scale radii of the thin- and thick-discs to be $R_{\text{t}} = 2.5$ kpc and $R_{\text{T}} = 2.0$ kpc, respectively, with a surface density ratio at $R_{\text{gal}} = 0$ of $\Sigma_{\text{T}}/\Sigma_{\text{t}} = 0.27$ based on the findings of Bland-Hawthorn & Gerhard (2016).

Since the Johnson et al. (2021) models run VICE with the SFH specified a priori, determining the gas supply requires an assumption regarding the SFE. Based on the findings of Kennicutt (1998), GCE models often adopt a single power-law relating the surface density of gas Σ_{gas} to the surface density of star formation $\dot{\Sigma}_{\star}$. Recent studies,

however, have revealed that the spatially resolved $\Sigma_{\text{gas}} - \dot{\Sigma}_{\star}$ relation is more nuanced than the integrated relation (de los Reyes & Kennicutt 2019; Ellison et al. 2021; Kennicutt & de los Reyes 2021). Some of the uncertainty regarding its details can be traced back to the ongoing debate about the CO-to-H₂ conversion factor (Kennicutt & Evans 2012; Liu, Gao & Greve 2015). Based on a compilation of the Bigiel et al. (2010) and Leroy et al. (2013) data shown in comparison to the theoretically motivated parametrizations of Krumholz et al. (2018, see their Fig. 2), Johnson et al. (2021) take a three-component power-law $\dot{\Sigma}_{\star} \propto \Sigma_{\text{gas}}^N$ with index N given by:

$$N = \begin{cases} 1.0 & (\Sigma_{\text{gas}} \geq 2 \times 10^7 \text{ M}_{\odot} \text{ kpc}^{-2}) \\ 3.6 & (5 \times 10^6 \text{ M}_{\odot} \text{ kpc}^{-2} \leq \Sigma_{\text{gas}} \leq 2 \times 10^7 \text{ M}_{\odot} \text{ kpc}^{-2}) \\ 1.7 & (\Sigma_{\text{gas}} \leq 5 \times 10^6 \text{ M}_{\odot} \text{ kpc}^{-2}). \end{cases} \quad (6)$$

The normalization is fixed by setting the SFE timescale $\tau_{\star} \equiv \Sigma_{\text{gas}} / \dot{\Sigma}_{\star}$ at surface densities where $N = 1$ to the value derived observationally for molecular gas, denoted by τ_{mol} . This value at the present day is taken to be $\tau_{\text{mol},0} = 2 \text{ Gyr}$ (Leroy et al. 2008, 2013), with a $t^{1/2}$ dependence on cosmic time based on the findings of Tacconi et al. (2018) studying the properties of molecular gas as a function of redshift.

With the yields adopted in the Johnson et al. (2021) models (see discussion at the beginning of § 2), considerable outflows are required in order to predict empirically plausible abundances. Weinberg et al. (2017) demonstrate analytically that the equilibrium abundance of some element in the gas phase is approximately determined by its yield and the mass loading factor $\eta \equiv \dot{\Sigma}_{\text{out}} / \dot{\Sigma}_{\star}$ with a small correction for the SFH. Johnson et al. (2021) select a scaling of η with R_{gal} such that the equilibrium abundance as a function of radius corresponds to a reasonable metallicity gradient within the Galaxy (see their Fig. 3 and discussion in their § 3.1). Nonetheless, yields and the strength of outflows are mutually degenerate parameters since they act as source and sink terms in computing enrichment rates. The absolute scale of nucleosynthetic yields is a topic of debate (see discussion in, e.g., Griffith et al. 2021), and some authors even neglect outflows

entirely, arguing that they do not significantly alter the chemical evolution of the Galaxy disc (e.g. Spitoni et al. 2019, 2021). We investigate the impact of simultaneous variations in our yields and the efficiency of outflows in § 4.2 below.

4 RESULTS

4.1 Evolution of a Fiducial Model

Our fiducial model adopts $y_{\text{N}}^{\text{CC}} = 3.6 \times 10^{-4}$ and our linear AGB star yield model (equation 3) with $\xi = 9 \times 10^{-4}$ along with the O and Fe yields of Johnson et al. (2021, see discussion in § 2). Although we presented this parametrization in comparison to the C11+C15 yields with $\xi = 3 \times 10^{-4}$ in Fig. 3, we find that a renormalization is necessary to reproduce the [N/O]-[O/H] relation as observed; we discuss this point further in § 4.2. To demonstrate the impact of stellar migration on N enrichment rates, we make use of the diffusion migration prescription in this section (see discussion in § 3).

In the left panel of Fig. 5, we plot the evolution of N and O abundances in the gas phase at five different Galactocentric radii. At early times, [O/H] is low and [N/O] reflects the ratio of the CCSN yields ($[\text{N/O}]_{\text{cc}} \approx -0.7$). Consequently, the tracks in each ring are similar. Once lower mass stars begin to evolve through an AGB phase, they enrich the interstellar medium (ISM) with N but negligible amounts of O, increasing [N/O]. At this point, the tracks in each ring separate from one another. This separation is a consequence of the metallicity gradient in [O/H] being established early in the Galaxy’s evolution. The radial gradient in our model arises out of a decrease in the equilibrium abundance of O with increasing radius. Produced on short timescales by CCSNe, O achieves equilibrium faster than elements produced by delayed nucleosynthetic sources (Weinberg et al. 2017). The ISM therefore reaches equilibrium in O soon after AGB stars begin producing N, after which [N/O] continues to increase at an approximately fixed [O/H] at all radii (see also Fig. 8 and associated discussion in § 4.4). The separation of these evolutionary tracks con-

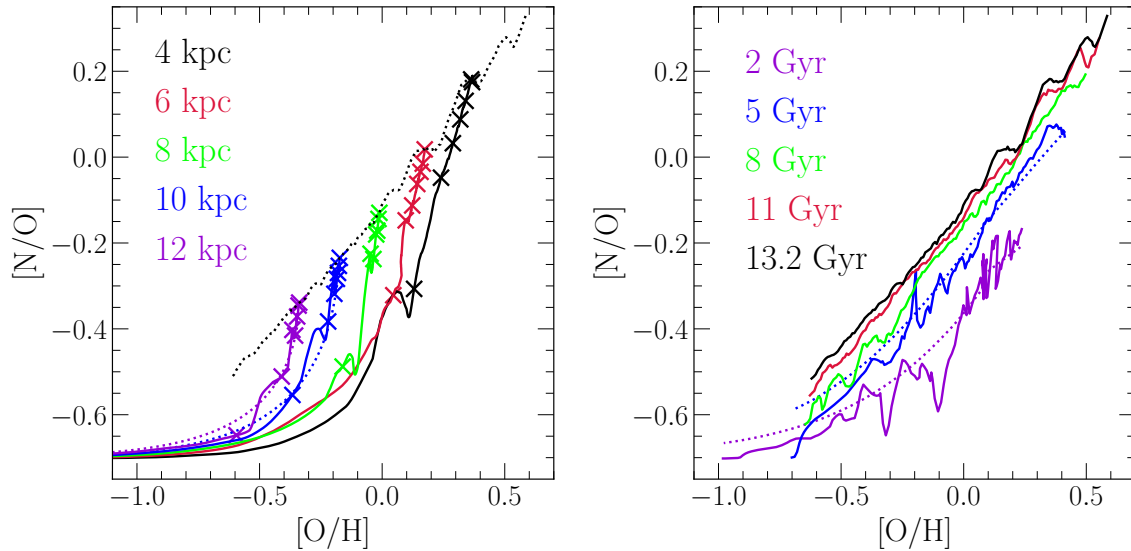


Figure 5. **Left:** The gas-phase $[N/O]$ - $[O/H]$ relation parametrized by time at fixed radius (solid coloured lines) in the fiducial model. X’s denote the abundances at $T = 2, 4, 6, 8, 10, 12$, and 13.2 Gyr (the present day) at all radii. The dotted black line is the same as the solid black line in the right panel. Coloured dotted lines mark the evolution of our model at $R_{\text{gal}} = 10$ and 12 kpc when we neglect the impact of stellar migration on enrichment rates (i.e. the “post-processing” migration prescription from Johnson et al. 2021; see discussion in § 3). **Right:** The gas-phase $[N/O]$ - $[O/H]$ relation parametrized by radius at various snapshots (solid coloured lines) in our fiducial model. Similar to the left panel, coloured dotted lines denote the resulting relation at $T = 2$ and 5 Gyr when we neglect stellar migration in computing enrichment rates.

tests the popular interpretation that the $[N/O]$ - $[O/H]$ relation arises as an evolutionary sequence, instead suggesting a superposition of evolutionary endpoints at equilibrium values set by different out-flow efficiencies. Similar arguments have been made regarding low $[\alpha/\text{Fe}]$ disc stars (e.g. Schönrich & Binney 2009; Nidever et al. 2014; Buck 2020; Sharma et al. 2021). In short, a given ring in our model does not evolve along the $[N/O]$ - $[O/H]$ relation, instead following this “rightward-then-upward” trajectory dictated by the timescales on which N and O achieve equilibrium.

Because there is a delay between a stellar population’s formation and N production from its AGB stars (~ 250 Myr in this model; see Fig. 4), stellar migration can in principle occur within this time interval. Although the bulk of migration occurs on longer timescales, this characteristic delay is comparable to the dynamical time of the Milky Way and is thus adequate for kinematic heating to at least begin. In zoom-in hydrodynamical simulations from the FIRE¹ suite (Hopkins et al. 2014), El-Badry et al. (2016) find that stars in a $M_{\star} \approx 10^{10.6} M_{\odot}$ galaxy can migrate $1 - 2$ kpc within 1 Gyr of their

formation. Consequently, N enrichment rates at fixed R_{gal} may differ significantly from their expected values given the SFH at that radius because stellar migration induced a deficit or surplus of N-producing AGB stars. These tracks can thus move vertically in the $[N/O]$ - $[O/H]$ plane in response to AGB stars moving between rings as the Galaxy evolves, producing the “jitter” in these evolutionary tracks. We demonstrate this effect by comparing the solid blue and purple lines to their dotted counterparts. These are the tracks we compute using the post-processing migration prescription which eliminates the impact of migration on enrichment rates (see discussion in § 3).

In the right panel of Fig. 5, we plot the gas-phase $[N/O]$ - $[O/H]$ relation predicted by the model at various snapshots. To obtain this, we simply take the N and O abundances in the ISM at a given output time for each $\delta R_{\text{gal}} = 100$ pc ring at $R_{\text{gal}} > 2$ kpc and plot them as a line. The relation is generally time-independent at $T \gtrsim 5$ Gyr. Although there is some slight evolution toward higher $[N/O]$, the total change in $[N/O]$ over this time interval is well within the intrinsic scatter derived observationally (see Fig. 1). Even at $T = 2$ Gyr, corresponding to $z \approx 2.6$, $[N/O]$ at fixed $[O/H]$ is only ~ 0.2 dex lower than its value

¹ <https://fire.northwestern.edu>

at the present day. Especially when considering the intrinsic scatter that would arise if we were to consider models with, e.g., different SFHs, this calculation supports previous arguments that the redshift evolution of the [N/O]-[O/H] relation is at most minimal (Vincenzo & Kobayashi 2018; Hayden-Pawson et al. 2021).

We again demonstrate the impact of stellar migration in the right panel of Fig. 6 by comparing the blue and purple solid lines to their dotted counterparts, which quantify the relation using the post-processing migration prescription. This indicates that the local jitter seen in the relation at a given time is a consequence of migration as discussed above. The mechanism by which stellar migration imposes these features in the [N/O]-[O/H] plane is qualitatively similar to what Johnson et al. (2021) find for SN Ia production of Fe. They found that the SN Ia rate in this model can vary by as much as a factor of ~ 3 at large radii ($R_{\text{gal}} \gtrsim 9$ kpc). When a deficit or surplus of SN Ia events is sustained for timescales comparable to the depletion time of the local ISM, the gas-phase abundance of Fe increases or decreases accordingly. As a consequence, some of the stellar populations that form during these events are Fe-poor enough to present as young stars ($\lesssim 6$ Gyr) with significantly super-solar $[\alpha/\text{Fe}]$ ratios, some of which are indeed observed in the solar neighbourhood with APOGEE² (Chiappini et al. 2015; Martig et al. 2015, 2016; Warfield et al. 2021). In the case of N, the effect is smaller ($\lesssim 0.1$ dex in [N/O]) because our model predicts N yields to be ejected from stellar populations ~ 5 times faster than Fe (Fig. 4). Consequently, there is less time for stellar migration to occur within the timescale of N production than there is within the timescale of Fe production.

² Although some of these stars appear to have mis-labeled ages (Jofré et al. 2016; Yong et al. 2016; Izzard et al. 2018; Silva Aguirre et al. 2018; Miglio et al. 2021), this phenomenon could explain an intrinsically young sub-component (Hekker & Johnson 2019; see discussion in §§ 3.1 and 3.4 of Johnson et al. 2021).

4.2 Comparison to Observed Gas-Phase Trends

We use the Dopita et al. (2016) [N/O]-[O/H] relation, which is inferred by fitting local stars and HII regions spanning a wide range of [O/H], as our observational benchmark. As seen in Fig. 1, the Pilyugin et al. (2010) calibration leads to a steeper relation at high [O/H]. Following Vincenzo et al. (2021), we adopt the Dopita et al. (2016) relation because it agrees well with the trends found for APOGEE disc stars and with results for MaNGA galaxies (Belfiore et al. 2017, see Fig. 1). None the less, uncertainties in the observed trends remain, and adopting a significantly different relation would lead to different conclusions about the metallicity dependence of N yields. To make the comparison between different yield models more clear, we neglect the impact of stellar migration on enrichment rates and make use of the post-processing migration prescription in this section (see discussion in § 3).

In the left panel of Fig. 6, we compare our model predictions with each of the AGB star yield tables predicted from stellar evolution models (see Fig. 3 and discussion in § 2.2). With no changes to the Johnson et al. (2021) parametrization of the model, all four AGB star yield tables fail to reproduce the [N/O]-[O/H] relation as observed. The C11+C15 and V13 yields are able to reproduce the qualitative trend, but with an incorrect normalization. The K10 and KL16+K18 yields, on the other hand, fail to reproduce the steadily sloped increase of [N/O] with [O/H]. The inverse dependence of [N/O] with [O/H] predicted by the K10 AGB star yields can be understood by the interaction between TDU and HBB (see discussion in § 2.2). Both effects are stronger at low metallicity, and since all of the K10 stars experiencing HBB also experience TDU (see their Table 1), such a result is unsurprising. This is also true for the KL16+K18 yields, but that model predicts a relatively flat [N/O]-[O/H] relation because of updated model inputs regarding opacity and mass loss and the impact this has on ^{14}N yields (see discussion in § 2.2).

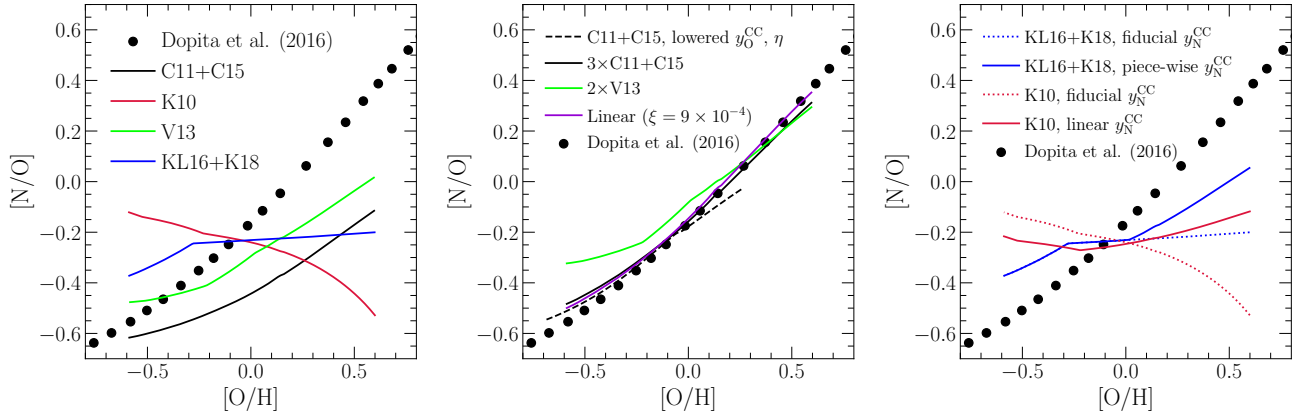


Figure 6. **Left:** The present-day gas-phase $[N/O]$ - $[O/H]$ relation predicted by our model with each of the four AGB star yield tables predicted by stellar evolution models discussed in § 2.2, colour-coded according to the legend. We include the [Dopita et al. \(2016\)](#) empirical relation as the observational benchmark. **Middle:** The same as the left panel, but for a case where we artificially amplify the **C11+C15** yields by a factor of 3 and the **V13** yields by a factor of 2. We show our fiducial model using the linear AGB star yields with a slope of $\xi = 9 \times 10^{-4}$ in a solid purple line. The black dashed line shows a model in which the **C11+C15** yields are unmodified but SN yields and the outflow mass loading factor are lowered by a factor of 3 (see discussion in § 4.2.1). **Right:** The same as the left panel, but comparing the predictions made by the **K10** and **KL16+K18** yields with our fiducial value of y_N^{CC} (dotted lines, same as left-hand panel) to those with alternate forms of y_N^{CC} (solid lines; see equations 7 and 8 and discussion in § 4.2). We show all predictions with our post-processing migration prescription (see discussion in § 3).

4.2.1 Variations in SN Yields and Outflows

In order to successfully reproduce the observations with the **C11+C15** and **V13** yields, we find that we must artificially amplify them by factors of ~ 3 and ~ 2 , respectively. We illustrate the results of these modified yield models and for our fiducial linear model with a normalization of $\xi = 9 \times 10^{-4}$ in the middle panel of Fig. 6. Although the **V13** model predicts an $[N/O]$ - $[O/H]$ relation that is slightly shallower than the [Dopita et al. \(2016\)](#) data, the predictions are reasonably within the scatter seen in Fig. 1.

As an alternative to amplifying the **C11+C15** and **V13** yields, we find good agreement with the observed relation if we instead lower our SN yields of N and O. The equilibrium $[N/O]$ ratio is largely determined by the IMF-averaged N and O yields, so lowering y_O^{CC} has much the same effect on $[N/O]$ as raising the AGB star N yields. However, lowering y_O^{CC} while holding other GCE parameters fixed reduces the O abundance at each Galactocentric radius, so the model no longer reproduces the normalization of the observed Galactic $[O/H]$ gradient. We can mostly repair this second problem by also reducing the outflow mass loading efficiencies assumed in the model. For a constant SFR, the equilibrium O mass fraction

is $Z_O = y_O^{CC} / (1 + \eta - r)$ where $r \approx 0.4$ is the recycling factor ([Weinberg et al. 2017](#)). The black dashed curve in the middle panel of Fig. 6 illustrates a model with the unmodified **C11+C15** yields, y_O^{CC} lowered by a factor of 3, and η lowered by a factor of 3 at all radii. This model exhibits good agreement with the [Dopita et al. \(2016\)](#) trend because N and O yields are both below those of the fiducial model by the same factor. The model reaches solar $[O/H]$ at the solar radius, but it predicts lower $[O/H]$ in the inner galaxy (i.e. a reduced x-axis span in Fig. 6) because lowering η by a factor of 3 does not lower $1 + \eta - r$ by a factor of 3.

Lowering our SN yields by a factor of 2–3 is plausible if a substantial fraction of massive stars collapse directly to black holes as opposed to exploding as SNe at the ends of their lives. Our IMF-averaged massive star yields (see discussion in § 2.1) are based on a [Kroupa \(2001\)](#) IMF combined with SN nucleosynthesis models in which most $M > 8 M_\odot$ stars explode as a CCSN (e.g. [Woosley & Weaver 1995](#); [Chieffi & Limongi 2004, 2013](#); [Limongi & Chieffi 2018](#); [Nomoto et al. 2013](#)). However, many massive stars may collapse to form black holes without a SN ([Pejcha & Thompson 2015](#); [Gerke, Kochanek & Stanek 2015](#); [Sukhbold et al. 2016](#); [Ertl et al. 2016](#); [Adams et al. 2017](#); [Basinger et al. 2021](#); [Neustadt et al. 2021](#)).

With the explosion landscape predicted by their W18 neutrino-driven engine, the CCSN models of [Sukhbold et al. \(2016\)](#) predict $y_{\text{O}}^{\text{CC}} = 0.0056$ ([Griffith et al. 2021](#)), nearly three times lower than our fiducial value.³ Extensive black hole formation would also lower $y_{\text{Fe}}^{\text{CC}}$, and a lower normalization of $y_{\text{Fe}}^{\text{Ia}}$ may be compatible with observational constraints on SN Ia rates (see discussion in § 2).

Another alternative is to retain high y_{O}^{CC} but assume that Galactic winds preferentially remove SN products relative to AGB products. In particular, [Vincenzo et al. \(2016a\)](#) are able to reproduce the [N/O]-[O/H] relation in chemical evolution models with the [V13](#) yields by implementing a differential wind in which outflows remove O but not N from the star forming gas reservoir. We find similar results for the [V13](#) yields if we simply add a portion of the SN products (both CCSN and SN Ia) directly to the outflow, which is otherwise composed of swept up ambient ISM with the same abundance ratios but reduced η . If SNe are the sources of outflow-driving winds but AGB stars do not significantly contribute, it would be reasonable to expect some portion of the SN ejecta to be swept up by the wind; recent theoretical ([Christensen et al. 2018](#)) and observational arguments ([Chisholm, Tremonti & Leitherer 2018](#)) indeed suggest such a scenario.

4.2.2 Metallicity-Dependent CCSN Yields of N

While the issue for the [C11+C15](#) and [V13](#) yields is one of normalization, our models with the [K10](#) or [KL16+K18](#) AGB yields predict a qualitatively incorrect trend of [N/O] with [O/H]. This discrepancy cannot be repaired by changing y_{O}^{CC} . However, the N yield is the sum of CCSN and AGB star contributions, so it is reasonable to ask if plausible changes to the metallicity dependence of the CCSN yield can compensate. Motivated by the observed [N/O] plateau at low metallicity and by the predictions of rotating massive star models, we

have thus far assumed a metallicity independent $y_{\text{N}}^{\text{CC}} = 3.6 \times 10^{-4}$. However, the non-rotating CCSN models of [Nomoto et al. \(2013\)](#) suggest that y_{N}^{CC} may increase at super-solar metallicity (see Fig. 3). We therefore construct the following parametrization for use with the [KL16+K18](#) AGB star yields:

$$y_{\text{N}}^{\text{CC}} = (3.6 \times 10^{-4}) \max\left(1, \frac{Z}{Z_{\odot}}\right). \quad (7)$$

Using this yield combination in our GCE model produces the blue solid curve in the right panel of Fig. 6. While agreement with the [Dopita et al. \(2016\)](#) trend is somewhat improved, this model is still far from the empirical [N/O]-[O/H] relation. Achieving agreement while using the [KL16+K18](#) yields would require still higher CCSN yields at $Z > Z_{\odot}$ and somewhat lower CCSN yields at $Z < Z_{\odot}$.

Because the [K10](#) AGB model predicts a high N yield at low metallicity (Fig. 4), we combine them with the predicted yields of the non-rotating CCSN models of [Limongi & Chieffi \(2018\)](#), which we approximate as

$$y_{\text{N}}^{\text{CC}} = (3.6 \times 10^{-4}) \left(\frac{Z}{Z_{\odot}}\right) \quad (8)$$

(diagonal dotted line in Fig. 2, left). This combination produces the red solid line in the right panel of Fig. 6. Although the discrepancy with the [Dopita et al. \(2016\)](#) trend is somewhat reduced, this model still underpredicts [N/O] at $Z > Z_{\odot}$ and overpredicts [N/O] at $Z < Z_{\odot}$ by large margins.

We conclude that the metallicity dependence of N yields predicted by the [KL16+K18](#) and, especially, [K10](#) AGB models are empirically untenable unless massive star yields are far from theoretical expectations. The implications for stellar astrophysics are uncertain, but inspection of Fig. 3 suggests that the problem of these AGB star models originates in the coexistence of TDU and HBB over a substantial mass range ($M \gtrsim 4 M_{\odot}$): in both models, every star that experiences HBB also experiences TDU.

³ This can be calculated with VICE using the `vice.yields.ccsne.fractional` function, designed to compute values of y_{X}^{CC} for various elements under a variety of assumptions.

4.2.3 Summary

The gas-phase $[N/O]$ - $[O/H]$ relation in the Milky Way disc predicted by our GCE model agrees well with the [Dopita et al. \(2016\)](#) trend characterizing local stars and HII regions if we assume a metallicity-independent $y_N^{CC} = 3.6 \times 10^{-4}$ as suggested by rotating massive star models ([Limongi & Chieffi 2018](#)) and the linear model of AGB yields in equation (3) with $\xi = 9 \times 10^{-4}$. Reproducing the empirical trend with the [C11+C15 \(V13\)](#) AGB yields requires renormalizing those yields by a factor of 3 (2) or, alternatively, lowering the assumed CCSN O yield y_O^{CC} by the same factor. Lowering y_O^{CC} could be physically justified if a large fraction of massive stars collapse to black holes without producing CCSNe, and it may be empirically tenable in a model where y_{Fe}^{CC} , y_{Fe}^{Ia} , and outflow mass loading efficiencies η are all lowered by a similar factor. The metallicity dependence of AGB star yields predicted by the [K10](#) or [KL16+K18](#) models, flat or even declining with increasing metallicity, is difficult to reconcile with the observed $[N/O]$ - $[O/H]$ trend.

4.3 Metallicity Dependence vs. Age Dependence

Despite predicting a different mass dependence for y_N^{AGB} (see Fig. 3), the renormalized [C11+C15](#) and [V13](#) yield models both reproduce the slope of the $[N/O]$ - $[O/H]$ relation reasonably well. This result suggests that the metallicity dependence plays a much more important role than the DTD in establishing this correlation. To investigate this point further, we consider two variants of our fiducial model: one with the dependence on stellar age (or, equivalently, stellar mass) removed from the enrichment rate calculations, and the other with the metallicity dependence removed. To make this comparison more straightforward, we use the post-processing migration model (see discussion in § 3).

To remove the age dependence, we simply eject the AGB star yields alongside the CCSN yield instantaneously after a single stellar population forms. We pre-compute the N yields from all AGB stars associated with a 13.2 Gyr old stellar population as a function of

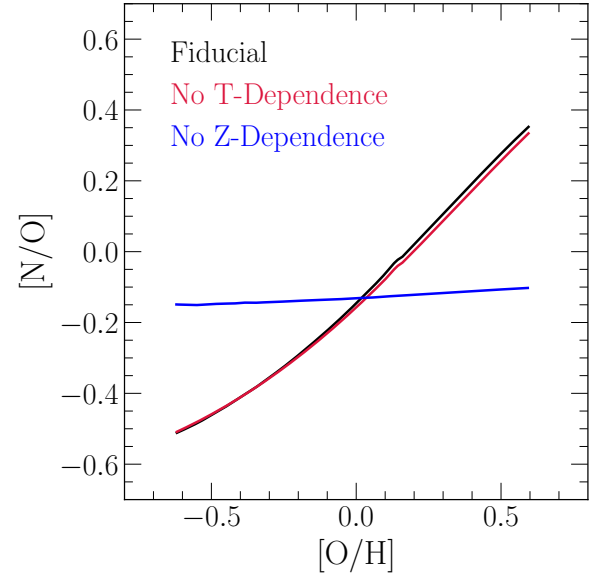


Figure 7. A comparison between our fiducial model with post-processing migration (black) and variations with the time dependence (red) and metallicity dependence removed (blue). To remove the time dependence, we pre-compute the AGB star yields of N from 13.2 Gyr old stellar populations as a function of metallicity as in the right panel of Fig. 4, then incorporate this into the prompt CCSN yields and set the delayed AGB star contribution to zero. To remove the metallicity dependence, we evaluate the yields at our assumed solar metallicity of $Z = 0.014$ at all timesteps.

progenitor metallicity in a similar fashion as in the right panel of Fig. 4; we retain the linear yield model with a normalization of $\xi = 9 \times 10^{-4}$ in this section. Since VICE works from IMF-averaged CCSN yields assumed to be injected instantaneously following a single stellar population’s formation (see discussion in § 2.1), we make use of the software’s capability to let the user specify functional forms for nucleosynthetic yields and simply add this N yield to y_N^{CC} and set y_N^{AGB} to zero. In this model, y_N^{CC} inherits a metallicity dependence from the AGB star yields and has the exact shape of the purple curve in the right hand panel of Fig. 4. To remove the metallicity dependence, the procedure is much simpler: we simply evaluate y_N^{AGB} at our assumed solar metallicity of $Z = 0.014$ at all timesteps regardless of that which is predicted for a stellar population. In this variation, AGB star production still occurs on a DTD inherited from the stellar mass-lifetime relation ([Larson 1974](#)) and the mass dependence of the linear yield model.

We illustrate these model predictions in Fig. 7. The $[N/O]$ - $[O/H]$ relation from the model with no age dependence is nearly

identical to the prediction found in our fiducial model, while the prediction with no metallicity dependence is considerably different. This result is rather unsurprising given the short characteristic timescales of N production (~ 250 Myr, see the middle panel of Fig. 4). Mathematically, there is little difference in the enrichment rates if all of a stellar population's N is produced immediately as opposed to from a prompt, sharply declining DTD. The metallicity dependence, however, is paramount to the $[\text{N/O}]-[\text{O/H}]$ relation, which is expected given the results in Fig. 3 and consistent with previous arguments that the increase in $[\text{N/O}]$ at high $[\text{O/H}]$ is a consequence of secondary N production (Vila-Costas & Edmunds 1993; van Zee et al. 1998; Henry & Worthey 1999; Pérez-Montero & Contini 2009; Berg et al. 2012; Pilyugin et al. 2012; Andrews & Martini 2013; Hayden-Pawson et al. 2021). Fig. 7 implies that the gas-phase $[\text{N/O}]-[\text{O/H}]$ relation offers little if any constraining power over the mass dependence of N yields from AGB stars.

4.4 Comparison to Stellar Abundances in the Milky Way Disc

Although N abundances are typically measured in the gas-phase in external galaxies, APOGEE (Majewski et al. 2017) has measured N abundances in large stellar samples spanning many regions of the Milky Way. By additionally making use of these data, we can investigate trends with stellar age and $[\text{O/Fe}]$ at fixed metallicity. Before comparing the predictions of GCE models to N abundances derived from spectra of red giant samples such as APOGEE, it is essential to adjust the measurements for internal processes known to alter the surface compositions of stars because GCE models predict the birth abundances. During the main sequence lifetime of $M \gtrsim 1.3 M_{\odot}$ stars, the CNO cycle processes much of the C and O nuclei in the core into ^{14}N . When the star evolves off the main sequence, this N-rich material is mixed with the outer convective layers, increasing the N abundance in the photosphere (Gilroy 1989; Korn et al. 2007; Lind et al. 2008; Souto et al. 2018, 2019). Using MESA stellar evolution models (Paxton et al. 2011, 2013, 2015, 2018) with standard

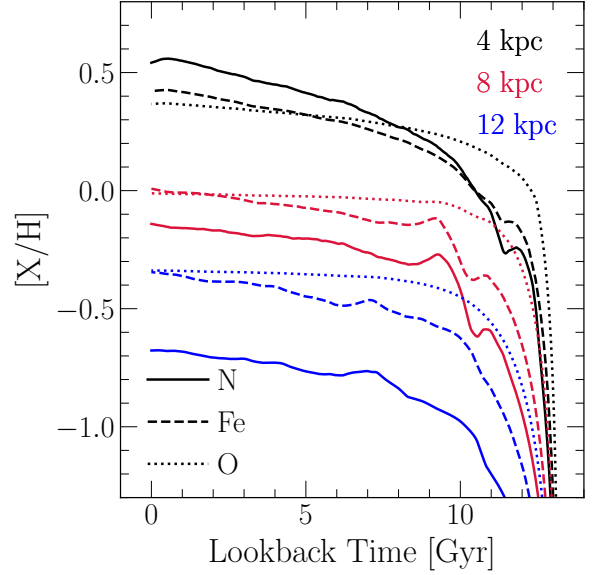


Figure 8. $[\text{N/H}]$ (solid), $[\text{Fe/H}]$ (dashed), and $[\text{O/H}]$ (dotted) in the gas-phase as a function of lookback time in the fiducial model at $R_{\text{gal}} = 4$ (black), 8 (red), and 12 kpc (blue).

mixing prescriptions, Vincenzo et al. (2021) developed a recipe to approximate the birth abundances of C, N, and O and applied it to the sample of APOGEE/Kepler red giants with asteroseismic mass measurements from Miglio et al. (2021). They found good agreement between the mean trend of $[\text{N/O}]$ with $[\text{O/H}]$ for APOGEE disc stars and the Dopita et al. (2016) trend. Since our fiducial model reproduces the Dopita et al. (2016) trend over most of its history (Fig. 5), it should also reproduce the $[\text{N/O}]-[\text{O/H}]$ relation for APOGEE disc stars.

Our models predict a correlation between N and Fe abundances in the gas-phase which turns out to be important to understanding how the model predictions compare to stellar abundances. In Fig. 8, we plot the evolution of $[\text{N/H}]$, $[\text{O/H}]$, and $[\text{Fe/H}]$ in the ISM at $R_{\text{gal}} = 4$, 8, and 12 kpc in our fiducial model with diffusion migration (see discussion in § 3). $[\text{N/H}]$ is more correlated with $[\text{Fe/H}]$ than $[\text{O/H}]$ at all radii, and the relation persists up to lookback times of ~ 10 Gyr. This arises in part because N and Fe are both produced in significant quantities by delayed enrichment sources while O is produced almost entirely on short timescales by CCSNe (see discussion in § 2). Although the production timescale of N from single stellar populations

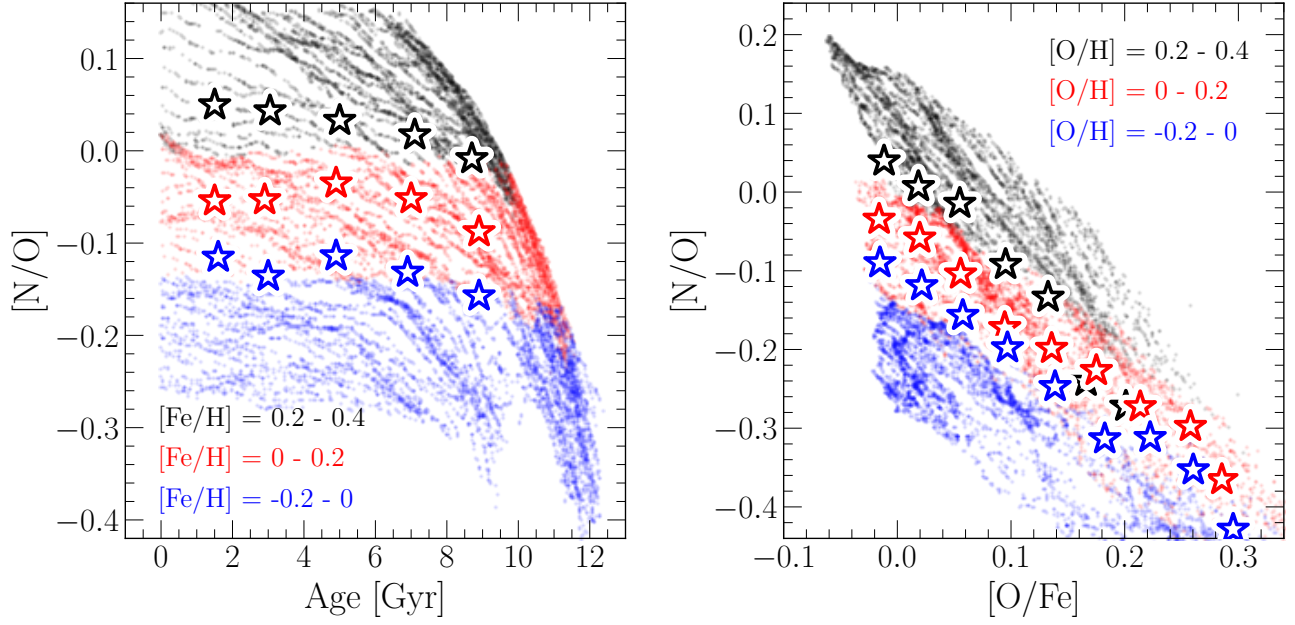


Figure 9. **Left:** $[N/O]$ as a function of stellar age for 5000 stars randomly sampled from our model stellar populations in three bins of $[Fe/H]$ (coloured points). Stars quantify the median trend in $[N/O]$ with age using N abundances corrected for internal mixing processes reported by Vincenzo et al. (2021) in the same bins of $[Fe/H]$. **Right:** The same as the left panel, but instead showing $[N/O]$ as a function of $[O/Fe]$ in bins of $[O/H]$.

is short (see discussion in § 2.3), metallicity dependent yields require more abundant species such as O to be produced and reach an equilibrium before N yields stabilize. When many stellar populations are present, the bulk of the N production will thus always follow the bulk production of more abundant species; this is qualitatively similar to what Johnson & Weinberg (2020) found regarding the production timescales of Sr and Fe. As a consequence of both its slight delay and its metallicity-dependent yields, N reaches its equilibrium abundance on timescales similar to Fe rather than O. Due to the prompt and metallicity independent nature of O enrichment, $[O/H]$ is near equilibrium as far back as ~ 10 Gyr ago while $[N/H]$ and $[Fe/H]$ are not.

Combining the Vincenzo et al. (2021) $[N/O]$ ratios with the APOGEE stellar ages taken from Miglio et al. (2021), we illustrate the $[N/O]$ -age relation in bins of $[Fe/H]$ as predicted by our model and measured by APOGEE in the left panel of Fig. 9. In good agreement with the observational measurements, the model predicts the $[N/O]$ -age relation to be relatively flat in bins of $[Fe/H]$. This arises as a consequence of the N-Fe correlation and the fast

approach to equilibrium in $[O/H]$ as discussed above (see Fig. 8). A bin in $[Fe/H]$ approximately corresponds to a bin in $[N/H]$, and by extension a bin in $[N/O]$ as well since $[O/H]$ is nearly constant at fixed radius up to ~ 10 Gyr ago.

In the right panel of Fig. 9, we compare our model predictions to the $[N/O]$ - $[O/Fe]$ relation at fixed $[O/H]$ reported by Vincenzo et al. (2021). The model correctly predicts a significant inverse relationship between $[N/O]$ and $[O/Fe]$. This is again a consequence of the N-Fe correlation demonstrated in Fig. 8: $[N/H]$ increases with $[Fe/H]$, so at fixed $[O/H]$, $[N/O]$ increases as $[O/Fe]$ decreases. This is another important success of our model. Vincenzo et al. (2021) demonstrate that high $[\alpha/Fe]$ and low $[\alpha/Fe]$ disc populations show a dichotomy in $[N/O]$. The Johnson et al. (2021) GCE model produces a broad but continuous $[\alpha/Fe]$ distribution rather than the bimodal distribution found in previous works (e.g. Hayden et al. 2015), but Fig. 9 suggests that a model tuned to produce the $[\alpha/Fe]$ bimodality would also produce a dichotomy in $[N/O]$.

Quantitatively, our model slightly overpredicts $[N/O]$ in the lower metallicity bins in both panels of Fig. 9. In general, our model occu-

pies a noticeably wider range in [N/O] than do the [Vincenzo et al. \(2021\)](#) measurements at all ages and all [O/Fe]. This could be a sign that the AGB star yields of N in our fiducial model scale slightly too strongly with the total metallicity Z . Since our fiducial model assumes an exactly linear scaling of the N yield with Z (see equation 3), this suggests that perhaps a slightly sub-linear scaling would be more accurate, but only barely because the discrepancies in Fig. 9 are at the ~ 0.1 dex level.

Although we demonstrate in § 4.3 that the metallicity dependence of the yield plays the strongest role in establishing the [N/O]-[O/H] relation, the DTD plays a significant role in shaping the stellar abundances. In the model in which AGB nucleosynthetic yields are injected instantaneously along with CCSN products, the N-Fe correlation described above is no longer present. Instead, N approaches equilibrium on much faster timescales, resulting in it being much more correlated with O than Fe, and the resulting [N/O]-age relation is positively sloped at fixed [Fe/H]. This suggests that the DTD may play a minimal role in establishing gas-phase abundances but are important in shaping stellar abundances. We find additional discrepancies if we instead attribute the delayed N production to a metallicity-dependent SN Ia yield with the same $t^{-1.1}$ DTD that we adopt for Fe. The agreement in Fig. 9 suggests that our fiducial model has a fairly accurate separation of CCSN and AGB contributions and that the ~ 250 Myr characteristic delay for AGB enrichment predicted by this model is approximately correct – instantaneous N enrichment is too fast and $t^{-1.1}$ enrichment is too slow to match the observations.

4.5 The Sources of Scatter in the [N/O]-[O/H] Relation

[Schaefer et al. \(2020\)](#) demonstrate that intrinsic scatter in the gas-phase [N/O]-[O/H] relation is correlated with variations in the local SFE. This is expected from one-zone GCE models (e.g. [Mollá et al. 2006](#); [Vincenzo et al. 2016a](#)). Although we have demonstrated in § 4.1 that the impact of stellar migration on enrichment rates is small, it could none the less contribute additional scatter in the ob-

served [N/O]-[O/H] relation. Our models, taking into account the effects of migration on the enrichment rates while allowing full control over the SFH and the SFE through VICE, are an ideal tool with which to address this question.

To this end, we construct two variants of our fiducial model. While the fiducial model specifies the SFH *a priori* and lets VICE compute the infall history $\dot{\Sigma}_{\text{in}}$, here we specify the infall history as a function of radius and time. As we will demonstrate below, the effects of dilution play an important role in driving variations in the [N/O]-[O/H] plane in these variants, and by specifying the infall history we have more control over the amount of dilution. In a similar fashion as in our fiducial model, we normalize $\dot{\Sigma}_{\text{in}}$ such that a stellar mass consistent with that reported by [Licquia & Newman \(2015\)](#) arises from the simulation. All other evolutionary parameters are the same as in the [Johnson et al. \(2021\)](#) fiducial model (see discussion in § 3 for details).

In the first variant, the SFE exhibits 50% sinusoidal oscillations with time over 2 Gyr periods:

$$\tau_{\star}(R_{\text{gal}}, t) = \tau_{\star, \text{J21}}(R_{\text{gal}}, t) \left(1 + 0.5 \sin \left(\frac{2\pi t}{2 \text{ Gyr}} \right) \right). \quad (9)$$

The infall rate is constant in each ring with a value determined by normalizing to the present day stellar mass and stellar surface density gradient of the Milky Way. Our choice of a 50% amplitude is comparable to the observationally derived scatter in molecular gas depletion times according to multiple measurement methods (see Figs. 4 and 5 of [Tacconi et al. 2018](#) and references therein). Furthermore, variations in the SFE are of similar magnitude in h277, the galaxy from which our model’s migration history is drawn. In the second variant, the SFE is constant and the infall rate oscillates with a 75% amplitude about its value in the first variant:

$$\dot{\Sigma}_{\text{in}}(R_{\text{gal}}, t) = \langle \dot{\Sigma}_{\text{in}} \rangle \left(1 + 0.75 \sin \left(\frac{2\pi t}{2 \text{ Gyr}} \right) \right). \quad (10)$$

The amplitude of 75% is chosen such that the ensuing variability in the SFR is of similar magnitude between the two models ($\sim 40\%$). We additionally run a model in which neither the accretion rate nor the SFE oscillate, replacing the fiducial model used in previous

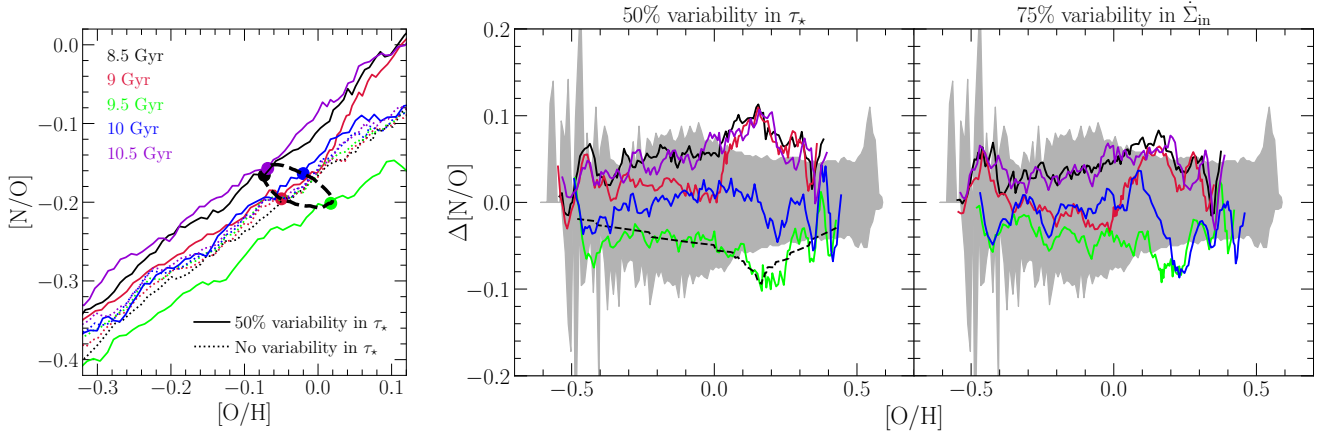


Figure 10. **Left:** One cycle of oscillations in the $[N/O]$ - $[O/H]$ relation at high $[O/H]$ induced by sinusoidal variability in τ_\star with an amplitude of 0.5 (solid coloured lines; see equation 9). Dotted lines show the $[N/O]$ - $[O/H]$ relation at the same five snapshots in the fiducial model with no variability in τ_\star . We use the diffusion migration prescription in both cases (see discussion in § 3). The black dashed line shows the time evolution of the abundances at $R_{\text{gal}} = 8$ kpc, the approximate Galactocentric radius of the sun, with the times of each of the five snapshots marked by a coloured point. **Middle and Right:** For the same five snapshots in the left hand panel, the deviation in $[N/O]$ at fixed $[O/H]$ relative to the fiducial model for the case with sinusoidal variability in τ_\star at an amplitude of 0.5 (middle; see equation 9) and with sinusoidal variability in $\dot{\Sigma}_\star$ at an amplitude of 0.75 (right; see equation 10). The shaded regions in both panels quantify the width of the $[N/O]$ distribution in $10^{10.5} - 10^{11} M_\odot$ galaxies in MaNGA taken from Schaefer et al. (2020). In bins of $[O/H]$, we place the median $[N/O]$ at $\Delta[N/O] = 0$, and the lower (upper) envelope denotes the 16th (84th) percentile of the $[N/O]$ distribution. The black dashed line in the middle panel denotes the same quantity as the corresponding solid green line but computed from the post-processing migration prescription, which neglects the impact of stellar migration in computing enrichment rates (see discussion in § 3).

sections with this constant infall model. Otherwise, the evolutionary differences beyond simple oscillations complicate the comparison.

These variant models characterize evolutionary pathways in which star formation is more episodic than previously explored. In a real galaxy, variability in the SFE and the SFR is likely non-sinusoidal and not with constant amplitude. A sample of galaxies will have different amplitudes and be seen at different phases in their variability, and the impact of this on their N and O abundances will present as intrinsic scatter in a sufficiently large sample. By comparing models with and without reasonable amounts of variability in these quantities while taking into account radial migration, we can assess which quantities impact abundances more strongly and are thus the more likely causes of intrinsic scatter in the observed $[N/O]$ - $[O/H]$ relation.

In the left panel of Fig. 10, we plot the predicted gas-phase $[N/O]$ - $[O/H]$ relation for five snapshots covering one cycle of fluctuations induced by variability in τ_\star according to equation (9). This model predicts a ~ 0.1 dex dynamic range in $[N/O]$ at fixed $[O/H]$, whereas the constant model with no variability in τ_\star predicts the relation to be quite steady over this time interval. This suggests that stellar migra-

tion, present in both the constant model and this oscillatory variant, does not induce significant variability in the $[N/O]$ - $[O/H]$ plane; however, we demonstrate below that its effects are none the less non-negligible. The minimal impact of stellar migration traces back to the timescales of N production from single stellar populations (see Fig. 4 and discussion in § 2.3): with most N production occurring within ~ 250 Myr of a stellar population’s formation, most stars will not migrate far from their birth radius by the time they produce most of their N, and the resulting impact on abundances is small.

The behavior in the $[N/O]$ - $[O/H]$ plane predicted by the oscillatory SFE variant is driven by the tug-of-war between dilution and re-enrichment associated with oscillations in τ_\star . When star formation quickens, O production increases in proportion. The ISM abundance and consequently the N yields increase as well. Because of the slight but none the less finite delay-time of its production by AGB stars, the N enrichment rate lags slightly behind O. $[N/O]$ therefore decreases, and the ISM moves down and to the right in the $[N/O]$ - $[O/H]$ plane. When star formation eventually slows, O production again follows suit. The N enrichment rate, as before, lags slightly

behind, and $[\text{N}/\text{O}]$ increases; the ISM therefore moves up and to the left in the $[\text{N}/\text{O}]-[\text{O}/\text{H}]$ plane. The result is an anti-clockwise loop, which we illustrate for the solar circle with a black dashed line in the left panel of Fig. 10. The effect is generally larger in $[\text{O}/\text{H}]$ than in $[\text{N}/\text{O}]$ (~ 0.1 dex versus ~ 0.05 dex in this example) because dilution affects both $[\text{O}/\text{H}]$ and $[\text{N}/\text{H}]$ similarly. In the model with oscillations in $\dot{\Sigma}_{\text{in}}$, we find that qualitatively similar processes drive the evolution in abundances, but there are interesting differences in detail which we discuss below in the context of scatter in the observed trend.

In the middle and right panels of Fig. 10, we plot the scatter in the gas-phase $[\text{N}/\text{O}]-[\text{O}/\text{H}]$ relation inferred observationally by Schaefer et al. (2020). Using data from the MaNGA IFU survey (Bundy et al. 2015), they measure N and O abundances in 709,541 spaxels across 6,507 unique galaxies spanning $10^9 - 10^{11} M_{\odot}$ in stellar mass. Since our model is appropriate for Milky Way mass galaxies, we focus our comparison on the $M_{\star} = 10^{10.5} - 10^{11} M_{\odot}$ mass range (Licquia & Newman 2015), which cuts our sample sample to 197,787 individual N and O measurements from the MaNGA IFU spaxels. In narrow bins of $[\text{O}/\text{H}]$, we then compute the 16th, 50th, and 84th percentiles of the $[\text{N}/\text{O}]$ distribution. Placing the median $[\text{N}/\text{O}]$ at $\Delta[\text{N}/\text{O}] = 0$, the shaded regions above and below 0 in Fig. 10 denote the difference between 16th and 84th percentiles of the distribution in each $[\text{O}/\text{H}]$ bin.

We compare both of our oscillatory variants to the width of the $[\text{N}/\text{O}]$ distribution by over-plotting the difference in $[\text{N}/\text{O}]$ at fixed $[\text{O}/\text{H}]$ between our oscillatory models and their constant counterpart (i.e. the vertical offset between the solid and dotted lines in the left panel, and the equivalent thereof for the oscillatory $\dot{\Sigma}_{\text{in}}$ model). Both models produce offsets in $[\text{N}/\text{O}]$ at fixed $[\text{O}/\text{H}]$ which, as discussed above, arise as consequences of dilution, and the offsets are generally consistent with the width of the relation derived observationally by Schaefer et al. (2020). This supports their argument that variations in the local SFE can drive intrinsic scatter in the $[\text{N}/\text{O}]-[\text{O}/\text{H}]$ relation, but the effects of stellar migration are still non-negligible. We demonstrate this by comparing the green solid

line in the middle panel for the $T = 9.5$ Gyr snapshot to the black dashed line denoting the same quantity with post-processing migration (see discussion in § 3). While some of the features in $\Delta[\text{N}/\text{O}]$ as a function of $[\text{O}/\text{H}]$ can be attributed to the difference in GCE parameters, stellar migration affects $[\text{N}/\text{O}]$ ratios with an amplitude of ~ 0.05 dex (see also Fig. 5 and discussion in § 4.1). Although this is smaller than the impact of oscillations in either τ_{\star} or $\dot{\Sigma}_{\text{in}}$ (~ 0.1 dex), it is none the less significant compared to the width of the Schaefer et al. (2020) distributions, also at the $\lesssim 0.1$ dex level; this suggests that stellar migration is a subdominant but non-negligible source of scatter.

In general, variability in τ_{\star} impacts abundances more strongly than variability in $\dot{\Sigma}_{\text{in}}$. Fig. 10 shows similar changes in $[\text{N}/\text{O}]$ at fixed $[\text{O}/\text{H}]$ in both of our oscillatory variants, but it requires an amplitude of 75% in accretion rates to achieve the same $\Delta[\text{N}/\text{O}]$ as an amplitude of 50% in the SFE. This weaker impact arises out of an abundance response that is much more *along* the $[\text{N}/\text{O}]-[\text{O}/\text{H}]$ relation rather than *against* it as in the oscillatory τ_{\star} model (see the black dashed line in the left panel of Fig. 10). Both $[\text{O}/\text{H}]$ and $[\text{N}/\text{H}]$ vary with larger amplitudes in the variable infall model due to episodes of enhanced and suppressed accretion, but the effects of dilution on $[\text{N}/\text{H}]$ are amplified by the combination with metallicity dependent yields. As a result, $[\text{N}/\text{H}]$ varies with a larger amplitude than $[\text{O}/\text{H}]$, whereas the opposite is the case in the oscillatory τ_{\star} model. Consequently, $[\text{N}/\text{O}]$ increases rather than decreases with increasing $[\text{O}/\text{H}]$, and changes in $[\text{N}/\text{O}]$ at fixed $[\text{O}/\text{H}]$ are smaller. In the context of the observational results (Schaefer et al. 2020), this suggests that different $[\text{N}/\text{O}]$ ratios at fixed $[\text{O}/\text{H}]$ are less likely to reflect changes in the accretion rate and more likely to reflect variations in the internal properties of the star forming ISM, such as the thermal state or the pressure, quantities which here get folded into τ_{\star} .

4.6 The [N/O]-[O/H] Relation as an Equilibrium Phenomenon

As shown by [Weinberg et al. \(2017\)](#), under generic conditions the abundances of a one-zone GCE model with continuing gas infall evolve to an equilibrium, in which new metal production is balanced by dilution and by the loss of ISM metals to star formation and outflows. Fig. 5 shows that the [N/O] ratio at a given radius in our multi-zone GCE model approaches an equilibrium after $t \approx 5$ Gyr. Fig. 7 further shows that the [N/O]-[O/H] relation that emerges in our model is driven by the metallicity dependence of the N yield with the time-delay of AGB enrichment having minimal impact. This indicates that – for the purposes of computing the equilibrium abundance of N – the AGB star DTD can be neglected, assuming instantaneous production as in § 4.3. This allows analytic solutions to the [N/O] ratio that will arise at a given equilibrium [O/H].

For an exponential SFH $\dot{M}_\star \propto e^{-t/\tau_{\text{SFH}}}$ with instantaneous enrichment and recycling of an element X with IMF-averaged yield y_X , the equilibrium ISM mass fraction is

$$Z_{\text{eq},X} = \frac{y_X}{1 + \eta - r - \tau_\star/\tau_{\text{SFH}}}, \quad (11)$$

where (as before) $\eta = \dot{M}_{\text{out}}/\dot{M}_\star$, $r \approx 0.4$ is the recycling fraction, and $\tau_\star = M_{\text{gas}}/\dot{M}_\star$ is the SFE timescale (see section 3). When an element’s characteristic enrichment delay time is $\ll \tau_{\text{SFH}}$, the correction to equation (11) is very small ([Weinberg et al. 2017](#)), which suggests in addition to Fig. 7 that the AGB DTD of N enrichment can safely be neglected for these purposes. The [N/O] abundance ratio in equilibrium is then given by the ratio of the yields with the metallicity-dependent N yield evaluated at the equilibrium O abundance:

$$[\text{N/O}]_{\text{eq}} = \log_{10} \left(\frac{Z_{\text{N,eq}}/Z_{\text{O,eq}}}{Z_{\text{N},\odot}/Z_{\text{O},\odot}} \right) = \log_{10} \left(\frac{y_{\text{N}}(Z_{\text{O,eq}})/y_{\text{O}}^{\text{CC}}}{Z_{\text{N},\odot}/Z_{\text{O},\odot}} \right), \quad (12)$$

where $[X/H] = \log_{10}(Z_X/Z_{X,\odot})$ and $y_{\text{N}}(Z_{\text{O}})$ denotes the *total* IMF-averaged N yield (CCSN + AGB) at a given Z_{O} as in Fig. 4. We have spot-checked this formula for each of our AGB yield models taken from the literature against the left panel of Fig. 6 and found agreement to 0.02 dex or better.

Equation (12) can be used to predict the [N/O]-[O/H] relation for a given set of yields. Our fiducial AGB star yield (equation 3 with $\xi = 9 \times 10^{-4}$) gives an IMF-averaged yield of 9.3×10^{-4} at $Z = Z_{\odot}$ (Fig. 4; see also discussion in § 2). Applying this value (along with $y_{\text{O}}^{\text{CC}} = 0.015$, $Z_{\text{N},\odot} = 6.91 \times 10^{-4}$, and $Z_{\text{O},\odot} = 5.72 \times 10^{-3}$; [Asplund et al. 2009](#)) to equation (12) and separating the massive star contribution of $y_{\text{N}}^{\text{CC}} = 3.6 \times 10^{-4}$ for generality gives

$$10^{[\text{N/O}]_{\text{eq}}} = 10^{[\text{N/O}]_{\text{cc}}} + (0.513)10^{[\text{O/H}]_{\text{eq}}}, \quad (13)$$

where $[\text{N/O}]_{\text{cc}} = -0.7$ is our empirical CCSN “plateau” taken from Fig. 1, for which alternate values can be computed from choices regarding y_{N}^{CC} and y_{O}^{CC} (see discussion in § 2.1). One can also reverse engineer equation (12) to derive the IMF-averaged N yield required to match an empirical [N/O]-[O/H] relation given a value of y_{O}^{CC} .

This analysis sharpens the conventional understanding of why [N/O] is a useful metallicity indicator in external galaxies when [O/H] cannot be measured directly. The relation between [N/O] and [O/H] is driven by stellar astrophysics (i.e. yields) in a way that is relatively insensitive to the SFH, SFE, or other galactic-scale physics. Viewing the relation from this perspective also highlights where one should be cautious about using [N/O] as a metallicity indicator. First is in environments where the stellar yields could be substantially different, perhaps because of differences in the IMF or in stellar rotation. Second is in galaxies that may be far from equilibrium, e.g. because of recent bursts of star formation or mergers diluting the ISM. The sinusoidal star formation variations considered in § 4.5 perturb equilibrium enough to create ~ 0.1 dex excursions in [N/O] at fixed [O/H] (Fig. 10). Third is in galaxies that are too young to have reached equilibrium. Although the [N/O]-[O/H] relation of our fiducial model is within ~ 0.1 dex of its final value by $t = 5$ Gyr, it is lower at earlier times (Fig. 5). One-zone models with a metallicity-dependent N yield can be used to estimate whether high-redshift galaxies are likely to have reached equilibrium based on their SFRs, gas fractions, metallicities, and stellar masses. Dwarf galaxies often have low SFE and might therefore reach equi-

librium more slowly. However, these galaxies are also typically in the low-metallicity regime where $[N/O]$ is determined by the metallicity-independent primary yields anyway.

5 CONCLUSIONS

Building on the multi-zone GCE model of [Johnson et al. \(2021\)](#), which reproduces many observed features of the $[\alpha/Fe]$ - $[Fe/H]$ -age distribution of Milky Way disc stars, we have inferred empirical constraints on the stellar nucleosynthesis of N by comparing model predictions to observed gas-phase trends in external galaxies and stellar trends in the Milky Way disc. In our models, the gas-phase abundance at a given Galactocentric radius first evolves to higher $[O/H]$ at roughly constant $[N/O]$ because of primary (metallicity-independent) N production, then evolves upward in $[N/O]$ with slowly increasing $[O/H]$ because of secondary N production that increases with metallicity (Fig. 5). The $[N/O]$ - $[O/H]$ relation reaches an approximate equilibrium after $t = 5 - 8$ Gyr, consistent with previous arguments that this relation is largely redshift-independent ([Vincenzo & Kobayashi 2018](#); [Hayden-Pawson et al. 2021](#)). This $[N/O]$ - $[O/H]$ relation represents a superposition of evolutionary track endpoints rather than an evolutionary track itself, similar to some explanations of the low- α disc sequence in the Milky Way (e.g. [Schönrich & Binney 2009](#); [Nidever et al. 2014](#); [Buck 2020](#); [Sharma et al. 2021](#); [Johnson et al. 2021](#)).

As our principal observational benchmark, we take [Dopita et al.'s \(2016\)](#) characterization of observed gas-phase abundances in external galaxies (see Fig. 1). Using [Johnson et al.'s \(2021\)](#) CCSN oxygen yield of $y_O^{CC} = 0.015$, we obtain agreement with the [Dopita et al. \(2016\)](#) $[N/O]$ - $[O/H]$ relation if we assume a metallicity-independent massive star yield of $y_N^{CC} = 3.6 \times 10^{-4}$ and an AGB fractional N yield that is linear in stellar mass and metallicity (equation 3) with $\xi = 9 \times 10^{-4}$. This value of y_N^{CC} is consistent with the rotating massive star models of [Limongi & Chieffi \(2018\)](#), and we concur with previous arguments that rotating massive stars are required to

explain the $[N/O] \approx -0.7$ plateau observed at low metallicities (see Fig. 1; [Chiappini et al. 2003, 2005, 2006](#); [Kobayashi et al. 2011](#); [Prantzos et al. 2018](#); [Grisoni et al. 2021](#)). The AGB yield is similar in form but 3 times higher in amplitude than the models of [C11+C15](#).

With $y_O^{CC} = 0.015$ and $y_N^{CC} = 3.6 \times 10^{-4}$, the AGB N yields of [C11+C15](#) or [V13](#) must be amplified by factors of three and two, respectively, to achieve agreement with the [Dopita et al. \(2016\)](#) $[N/O]$ - $[O/H]$ relation (Fig. 6). However, as predicted abundance ratios depend primarily on yield ratios, we can also obtain agreement by using the [C11+C15](#) or [V13](#) yields and lowering y_O^{CC} and y_N^{CC} by the corresponding factor. Such a change could be physically justified if black hole formation is more extensive, or the IMF steeper, than implicitly assumed by the value of $y_O^{CC} = 0.015$ (see § 4.2.1 and [Griffith et al. 2021](#)). Other successful predictions of the [Johnson et al. \(2021\)](#) models, including the Galactic $[O/H]$ gradient that is one of its basic constraints, would be largely unchanged if y_{Fe}^{CC} , y_{Fe}^{Ia} , and outflow mass loading efficiencies η were all reduced by the same factor. Alternatively, one could retain a higher y_O^{CC} and y_N^{CC} but assume that Galactic winds preferentially eject CCSN products relative to AGB products, as suggested by [Vincenzo et al. \(2016a\)](#). The degeneracy between the overall scaling of yields and the magnitude of outflows is one of the key sources of uncertainty in GCE models.

In contrast to [C11+C15](#) and [V13](#), the AGB models of [K10](#) and [KL16+K18](#) predict IMF-averaged yields that are decreasing or approximately flat with increasing Z (Fig. 4). In our GCE models, these yields lead to clear disagreement with the [Dopita et al. \(2016\)](#) trend, even when we allow reasonable variations in the metallicity dependence of y_N^{CC} (Fig. 6). There are many uncertain physical effects in AGB stellar models, so it is difficult to pinpoint a single cause for this discrepancy. In general, the most efficient N production occurs when both TDU and HBB occur simultaneously because each replenishment of C and O isotopes from the stellar core by TDU adds new seed nuclei for HBB to process in ^{14}N via the CNO cycle ([Ventura et al. 2013](#)). The distinctive metallicity dependence

of the K10 and KL16+K18 yields traces back to the simultaneous occurrence of TDU and HBB over a substantial mass range at all metallicities (Fig. 3).

All of the AGB models predict that IMF-averaged N production is dominated by stars with $M > 2 M_{\odot}$ (Fig. 4). As a result, the delay time required to produce 50% of the AGB N is 250 Myr or less, shorter than the ~ 1 Gyr characteristic delay of Fe from SN Ia. The form of the $[N/O]$ - $[O/H]$ relation is driven by the metallicity dependence of N yields, not by the time delay of AGB production (Fig. 7), and it can be calculated accurately from simple equilibrium arguments under most circumstances (§ 4.6, equation 12).

Vincenzo et al. (2021) inferred the median $[N/O]$ - $[O/H]$ trend of Milky Way disc stars from APOGEE abundances corrected for mixing on the red giant branch using the asteroseismic mass measurements from Miglio et al. (2021). They found good agreement with the Dopita et al. (2016) trend, our observational benchmark, so our model is also consistent with their derived APOGEE trends. Our model also reproduces, at least approximately, two important findings of Vincenzo et al. (2021): $[N/O]$ exhibits little correlation with stellar age at fixed $[Fe/H]$ for ages $\lesssim 9$ Gyr, and $[N/O]$ declines linearly with increasing $[O/Fe]$ at fixed $[O/H]$ (Fig. 9). The match to these observations *does* depend on the AGB DTD, and it breaks down if we either make the AGB enrichment instantaneous or make it occur as slowly as SN Ia Fe production.

To investigate the sources of scatter in the $[N/O]$ - $[O/H]$ relation, we construct variants of our fiducial model that have $\sim 40\%$ sinusoidal oscillations in the SFR with a 2 Gyr period, induced by oscillations in either the SFE or the gas infall rate. The combined effects of dilution by pristine infall and metallicity-dependent N production lead to oscillations in the $[N/O]$ - $[O/H]$ relation comparable in magnitude to the scatter measured in MaNGA galaxies by Schaefer et al. (2020) (Fig. 10). We concur with their conclusion that variations in the SFE can plausibly explain most of the observed scatter. Johnson et al. (2021) find that stellar migration induces stochastic variations

in $[\alpha/Fe]$ enrichment because a stellar population can migrate from its birth radius before most of its SN Ia Fe production takes place. The same effect occurs for AGB N enrichment but to a lesser extent because the shorter production timescale (~ 250 Myr) leaves less time for migration. We find that migration leads to ~ 0.05 -dex scatter in $[N/O]$ at fixed $[O/H]$, which is smaller than the scatter measured by Schaefer et al. (2020) but not negligible.

Our findings illustrate the value and methodology of empirically constraining stellar yields by combining general theoretical expectations with GCE modelling and observational constraints. For the case of N, we have used the expectation that massive stars and AGB stars both contribute, with the AGB contribution moderately delayed in time. The metallicity dependence of the combined IMF-averaged yield is tightly constrained, and it is plausibly partitioned into a massive star yield that is independent of metallicity and an AGB yield that is linear in metallicity Z and progenitor mass M . The normalization of the yield is well-constrained *relative* to the IMF-averaged O yield. The DTD predicted by our fiducial model, in concert with the Johnson et al. (2021) GCE prescriptions, leads to good agreement with the $[N/O]$ -age and $[N/O]$ - $[O/Fe]$ trends for Milky Way disc stars (Fig. 9). As this approach is extended to increasing numbers of elements, the web of yield constraints and consistency tests will become steadily more powerful, providing valuable insights on stellar astrophysics, SN physics, and the history of our Galaxy.

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7 DATA AVAILABILITY

VICE is open-source software. The code which numerically integrates our GCE models is also publicly available and can be found in VICE's GitHub repository. The sample of star particles from h277 is available through VICE, which will download the files automatically the first time it needs them. The rest of the data from h277 can be accessed at <https://nbody.shop/data.html>. Observational data taken from the literature (e.g. [Dopita et al. 2016](#); [Schaefer et al. 2020](#); [Vincenzo et al. 2021](#)) will be provided upon request to the corresponding author.

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Appendices

A VICE

VICE¹ is an open-source PYTHON package designed to model chemical enrichment processes in galaxies with a generic, flexible model. With this paper, we mark the release of version 1.3.0, which presents a handful of new features:

(i) Users may select a mass-lifetime relation for stars from a list of several parametrized forms taken from the literature. Previously, only a single power-law was implemented, but this formulation underestimates lifetimes for stars with masses $\gtrsim 4M_{\odot}$; now, the options include the equations presented in:

- [Vincenzo et al. \(2016b\)](#)
- [Hurley, Pols & Tout \(2000\)](#)
- [Kodama & Arimoto \(1997\)](#)
- [Padovani & Matteucci \(1993\)](#)
- [Maeder & Meynet \(1989\)](#)
- [Larson \(1974\)](#) (default)

Generally, chemical evolution models make similar predictions with each of these different forms of the mass-lifetime relation since they are not considerably different from one another (see the section titled “Single Stellar Populations” under VICE’s science documentation for further discussion²). We select the [Larson \(1974\)](#) form as a default within VICE because it is representative of other forms and requires the lowest amount of computational overhead (aside from the single power-law option). This model is a metallicity-independent parabola in $\log \tau - \log m$ space for which we take updated coefficients from [Kobayashi \(2004\)](#) and [David et al. \(1990\)](#).

(ii) We have added two additional tables of AGB star

¹ Install (PyPI): <https://pypi.org/project/vice>

Documentation: <https://vice-astro.readthedocs.io>

Source Code: <https://github.com/giganano/VICE.git>

² https://vice-astro.readthedocs.io/en/latest/science_documentation/

yields sampled at various progenitor masses and metallicities: the [KL16+K18](#) and [V13](#) models presented in this paper are new to VICE (see discussion in § 2.2 for details).

(iii) We have built in the SN Ia yields presented in [Gronow et al. \(2021a,b\)](#). These tables include yields for double detonations of sub-Chandrasekhar mass carbon-oxygen white dwarfs at various progenitor metallicities.

Although VICE includes built-in SN and AGB star yield tables, users are not required to adopt any one of them for use in their chemical evolution models. Instead, it allows arbitrary functions of metallicity for both CCSN and SN Ia yields and functions of progenitor mass and metallicity for AGB star yields. It provides similar flexibility for additional parameters typically built into GCE models. VICE's backend is implemented entirely in ANSI/ISO C, providing it with the powerful computing speeds of a compiled library while retaining such scientific flexibility within the easy-to-use framework of PYTHON.

Requiring a Unix kernel, VICE supports Mac and Linux operating systems; Windows users should install and use VICE entirely within the Windows Subsystem for Linux. It can be installed via `python -m pip install vice`, after which `python -m vice --docs` and `python -m vice --tutorial` will launch a web browser to the documentation and to a jupyter notebook intended to familiarize first time users with VICE's API.