

RESEARCH IN PHYSICS COURSE REPORT

Evolve Schwarzschild geometry using ADM equations

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Abstract

We have briefly reviewed the 3+1 formalism in general relativity which is essential for numerically solving Einstein's equations. Based on a simple form of this formalism, i.e., ADM formalism, preliminary results on the numerical solution for a spherically symmetric vacuum spacetime, e.g., the Schwarzschild metric, are presented. From this, we lay the groundwork for further simulations in the future such as stellar collapses and binary black hole coalescences.

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List of Abbreviations

NR	Numerical Relativity
ADM	Arnowitt, Deser and Misner
BSSN	Baumgarte, Shapiro, Shibata and Nakamura

List of Symbols

g_{ab}, g^{ab}	spacetime metric and its inverse
γ_{ij}, γ^{ij}	spatial metric and its inverse
∂_t	time vector
N	lapse function
β	shift vector
Σ_t	hypersurface at time t
${}^{(4)}R, {}^{(4)}R$	Ricci tensor associated with g and the corresponding Ricci scalar
T	matter stress-energy tensor
E	matter energy density
p	matter momentum density
S	matter stress tensor
\mathcal{L}_m	Lie derivative along vector field m

Chapter 1

Introduction

1.1 Why we need numerical relativity?

General relativity is the theory that best describes modern cosmology, celestial bodies such as black holes, and gravitational waves. In general relativity, Einstein's field equation describes curved spacetime as the momentum and energy of matter.

However, these are ten nonlinear partial differential equations, whose complete solutions have not yet been solved analytically, except for the Friedmann–Lemaître–Robertson–Walker metric, etc. With advances in computer algorithms and physics, Einstein's equations began to be solved numerically, which is the beginning of numerical relativity.

1.2 Introduction to numerical relativity

In numerical relativity, Einstein's equations is separated into space part and time part, that is, 3+1 decomposition.

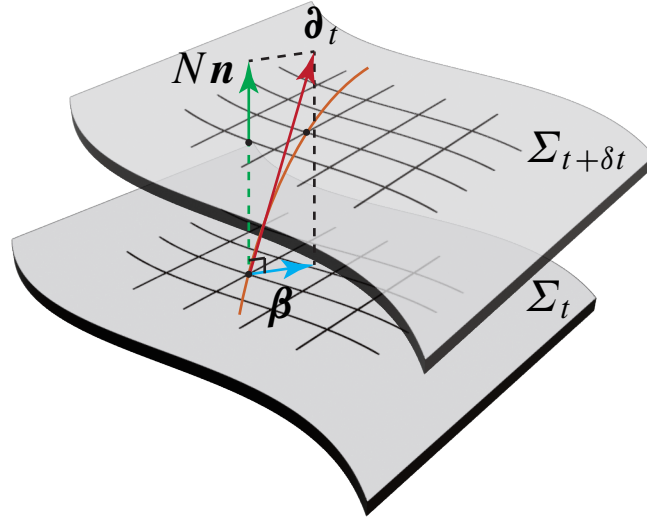


FIGURE 1.1: A foliation of the spacetime. There are time vector ∂_t , normal evolution vector $N\mathbf{n}$, and shift vector β satisfying $\partial_t = N\mathbf{n} + \beta$ between the two hypersurfaces $\Sigma_t, \Sigma_{t+\delta t}$.

From this, it is possible to calculate the next spatial metric from the initial condition of the space at a specific time. This is done in a similar way as in Newtonian classical mechanics, given the initial position and velocity of an object, the position and velocity at the next time can be calculated using the acceleration due to the force acting on the object. However, unlike in Newtonian mechanics, the initial value must satisfy certain specific conditions. Therefore, the calculation is made by first constructing the initial data, selecting the appropriate coordinates, and then evolving it from the numerical method.

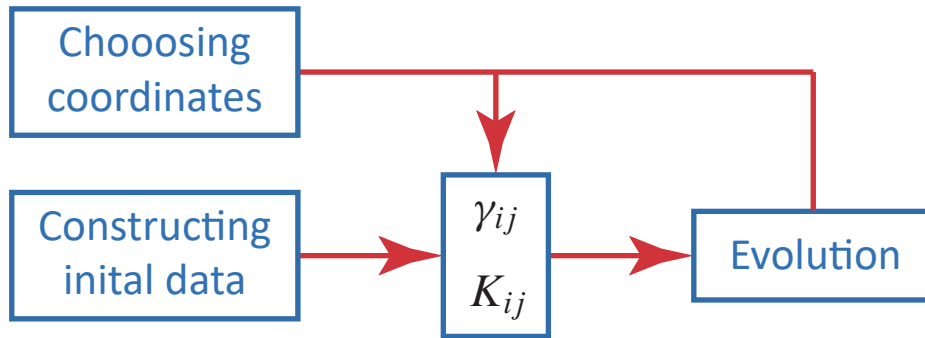


FIGURE 1.2: A brief schema of numerical relativity.

In this work, we try to numerically evolve the Schwarzschild metric from its initial conditions to understand the basic way of working with numerical relativity.

1.3 Notation and conventions

Throughout this report, we follow the “Landau-Lifshitz Spacelike Convention” (− + + +) as [1]. Also we will adopt a units for measurements in which both the gravitational constant G and the speed of light c are assigned the values of one.

We denote the dimension 4 spacetime metric by g_{ab} , the dimension 3 spatial metric by γ_{ij} . Also the dimension 4 objects associated with g_{ab} are denoted with a superscript ⁽⁴⁾ in front of the symbol, objects related to γ_{ij} carry no decorations.

Chapter 2

The ADM Equations

2.1 Introduction

In this chapter, we present how Einstein's equations can be decomposed into two constraints and two evolution equations. The development in this chapter follows [6].

2.2 The Einstein Equation

Einstein's equation is

$${}^{(4)}R - \frac{1}{2}{}^{(4)}Rg = 8\pi T. \quad (2.1)$$

We limit cosmological constant $\Lambda = 0$.

2.3 Constraint equations

Let's start with the *Gauss relation*

$$\gamma_\alpha^\mu \gamma_\beta^\nu \gamma_\rho^\gamma \gamma_\delta^\sigma {}^{(4)}R_{\sigma\mu\nu}^\rho = R_{\delta\alpha\beta}^\gamma + K_\alpha^\gamma K_{\delta\beta} - K_\beta^\gamma K_{\alpha\delta}. \quad (2.2)$$

We can obtain *scalar Gauss relation*

$${}^{(4)}R + 2{}^{(4)}R_{\mu\nu}n^\mu n^\nu = R + K^2 - K_{ij}K^{ij} \quad (2.3)$$

by contract the Gauss relation Eq. (2.2) on the indices γ and α and use $\gamma_\alpha^\mu \gamma_\rho^\alpha = \gamma_\rho^\mu = \delta_\rho^\mu + n^\mu n_\rho$, and take its trace with respect to γ .

From Eq. (2.1), after full projection perpendicular to hypersurface Σ_t , we get

$${}^{(4)}R(\mathbf{n}, \mathbf{n}) + \frac{1}{2}{}^{(4)}R = 8\pi T(\mathbf{n}, \mathbf{n}) =: 8\pi E \quad (2.4)$$

since $g(\mathbf{n}, \mathbf{n}) = -1$. By combining Eq. (2.3), (2.4), we get

$$R + K^2 - K_{ij}^{ij} = 16\pi E \quad (2.5)$$

which is called the *Hamiltonian constraint*.

Now let us project Eq. (2.1) onto Σ_t and normal \mathbf{n} ,

$${}^{(4)}R(\mathbf{n}, \vec{\gamma}(\cdot)) - \frac{1}{2}{}^{(4)}Rg(\mathbf{n}, \vec{\gamma}(\cdot)) = 8\pi T(\mathbf{n}, \vec{\gamma}(\cdot)). \quad (2.6)$$

We can use *Codazzi relation*

$$\gamma_\rho^\gamma n^\sigma \gamma_\alpha^\mu \gamma_\beta^\nu {}^{(4)}R_{\sigma\mu\nu}^\rho = D_\beta K_\alpha^\gamma - D_\alpha K_\beta^\gamma \quad (2.7)$$

to get *contracted Codazzi relation*

$$\gamma_\alpha^\mu n^\nu {}^{(4)}R_{\mu\nu} = D_\alpha K - D_\mu K_\alpha^\mu \quad (2.8)$$

by contracting the Eq. (2.7) on the indices α and γ .

Since $g(\mathbf{n}, \vec{\gamma}(\cdot))$ in Eq. (2.6) is equal to 0 and by introducing *matter momentum density* $\mathbf{p} := -T(\mathbf{n}, \vec{\gamma}(\cdot))$, we get

$$\mathbf{D} \cdot \vec{\mathbf{K}} - \mathbf{D} K = 8\pi \mathbf{p}, \quad (2.9)$$

or, in components,

$$D_j K_i^j - D_i K = 8\pi p_i \quad (2.10)$$

which is called the *momentum constraint*.

2.4 Evolution equations

Since we can write time vector in to sum of the normal evolution vector $\mathbf{m} := N\mathbf{n}$ and the shift vector $\boldsymbol{\beta}$,

$$\partial_t =: \mathbf{m} + \boldsymbol{\beta}, \quad (2.11)$$

we can write

$$\mathcal{L}_{\mathbf{m}} \mathbf{T} = \mathcal{L}_{\partial_t} \mathbf{T} - \mathcal{L} \mathbf{T}, \quad (2.12)$$

for \mathbf{T} be any tensor field tangent to Σ_t . Moreover, Lie derivative is simply obtained by taking the partial derivative of the vector components with respect to t . Therefore, Eq. (2.12) can be written as

$$\mathcal{L}_{\mathbf{m}} T_{j\cdots}^{i\cdots} = \left(\frac{\partial}{\partial t} - \mathcal{L}_{\boldsymbol{\beta}} \right) T_{j\cdots}^{i\cdots}. \quad (2.13)$$

By applying it to extrinsic curvature

$$\mathcal{L}_{\mathbf{m}} \boldsymbol{\gamma} = -2N \mathbf{K}, \quad (2.14)$$

it becomes

$$\left(\frac{\partial}{\partial t} - \mathcal{L}_{\boldsymbol{\beta}} \right) \gamma_{ij} = -2N K_{ij}, \quad (2.15)$$

which is called the *evolution equation for the spatial metric*.

If we applying the operator $\vec{\gamma}^*$ to the Einstein equation,

$$\vec{\gamma}^{*(4)} \mathbf{R} = 8\pi \left(\vec{\gamma}^* \mathbf{T} - \frac{1}{2} T \vec{\gamma}^* \mathbf{g} \right). \quad (2.16)$$

From the 3+1 decomposition of the Riemann tensor, we obtained

$$\vec{\gamma}^{*(4)} \mathbf{R} = -\frac{1}{N} \mathcal{L}_{\mathbf{m}} \mathbf{K} - \frac{1}{N} \mathbf{D} \mathbf{D} N + \mathbf{R} + \mathbf{K} \mathbf{K} - 2\mathbf{K} \cdot \vec{\mathbf{K}}. \quad (2.17)$$

Therefore

$$-\frac{1}{N} \mathcal{L}_{\mathbf{m}} \mathbf{K} - \frac{1}{N} \mathbf{D} \mathbf{D} N + \mathbf{R} + \mathbf{K} \mathbf{K} - 2\mathbf{K} \cdot \vec{\mathbf{K}} = 8\pi \left[\mathbf{S} - \frac{1}{2} (S - E) \boldsymbol{\gamma} \right], \quad (2.18)$$

where *matter stress tensor* $\mathbf{S} := \vec{\gamma}^* \mathbf{T}$.

The result from property that the Lie derivative along \mathbf{m} of any tensor field \mathbf{T} tangent to Σ_t is a tensor field tangent to Σ_t , Eq. (2.18) can be written as

$$\mathcal{L}_{\mathbf{m}} K_{ij} = -D_i D_j N + N \left\{ R_{ij} + K K_{ij} - 2K_{ik} K_j^k + 4\pi [(S - E) \gamma_{ij} - 2S_{ij}] \right\}. \quad (2.19)$$

From Eq. (2.11), we get *evolution equation for the extrinsic curvature*

$$\left(\frac{\partial}{\partial t} - \mathcal{L}_{\boldsymbol{\beta}}\right)K_{ij} = -D_i D_j N + N \left\{ R_{ij} + K K_{ij} - 2K_{ik} K_j^k + 4\pi[(S - E)\gamma_{ij} - 2S_{ij}] \right\}. \quad (2.20)$$

2.5 Summary

In this chapter, we obtained the *Hamiltonian constraint*

$$R + K^2 - K_{ij} K^{ij} = 16\pi E, \quad (2.21)$$

the *momentum constraint*

$$D_j K_i^j - D_i K = 8\pi p_i, \quad (2.22)$$

the *evolution equation for the spatial metric*

$$\left(\frac{\partial}{\partial t} - \mathcal{L}_{\boldsymbol{\beta}}\right)\gamma_{ij} = -2N K_{ij}, \quad (2.23)$$

and the *evolution equation for the extrinsic curvature*

$$\left(\frac{\partial}{\partial t} - \mathcal{L}_{\boldsymbol{\beta}}\right)K_{ij} = -D_i D_j N + N \left\{ R_{ij} + K K_{ij} - 2K_{ik} K_j^k + 4\pi[(S - E)\gamma_{ij} - 2S_{ij}] \right\}. \quad (2.24)$$

Chapter 3

Numerical simulation

3.1 Introduction

In this chapter, we describe the initial data set required prior to simulation, the numerical method for evolution, and the results.

3.2 Schwarzschild black hole

3.2.1 Types of black holes

Accordingly to the No-Hair theorem, all black holes solutions of the Einstein-Maxwell equation of electromagnetism in general relativity can be completely characterized by their observable classical parameters mass, electric charge and angular momentum.

The schwarzschild metric describes the spacetime geometry exterior to any spherical collapsing body. Kerr metric describes the geometry of empty spacetime around a rotating uncharged axially-symmetric black hole with quasi-spherical event horizon. There are also the Reissner-Nordström metric and the Kerr-Newman metric that describe charged black holes[8, 5]. In Table 3.1, it can be seen that the types of black holes are classified according to angular momentum and charge.

TABLE 3.1: Classifications of black holes.

	Non-rotating ($J = 0$)	Rotating ($J > 0$)
Uncharged ($Q = 0$)	Schwarzschild	Kerr
Charged ($Q \neq 0$)	Reissner-Nordström	Kerr-Newman

3.2.2 Isotropic coordinate

In this report, we choose the simplest form, the Schwarzschild black hole. The original form of the Schwarzschild metric is

$$ds^2 = -\left(1 - \frac{2M}{r}\right) dt^2 + \left(1 - \frac{2M}{r}\right)^{-1} dr^2 + r^2(d\theta^2 + \sin^2 \theta d\phi^2). \quad (3.1)$$

When r goes $2M$, g_{rr} diverges. However, this is just a coordinate singularity, like the problem that occurs at the north and south poles in the spherical coordinate system[7]. This can be solved by choosing another coordinate system, such as the Kruskal coordinate system.

We can avoid coordinate singularity at $r = 2M$ by adopting an isotropic coordinate system by substituting $r = \bar{r}(1 + M/2\bar{r})^2$, we get

$$ds^2 = -\left(\frac{1 - M/(2\bar{r})}{1 + M/(2\bar{r})}\right)^2 dt^2 + \left(1 + \frac{M}{2\bar{r}}\right)^4 (d\bar{r}^2 + \bar{r}^2 d\theta^2 + \bar{r}^2 \sin^2 \theta d\phi^2). \quad (3.2)$$

This coordinate system describes the area outside the event horizon $\bar{r} = M/2$. The reason for using this coordinate system is as follows.

1. A spatial metric is numerically valid in any space where $r > M/2$.
2. Since spatial metrics are flat, they can be replaced with Cartesian coordinates, which is more useful for numerical calculations[2].

So, in practice we use metric:

$$ds^2 = -\left(\frac{1 - M/(2r)}{1 + M/(2r)}\right)^2 dt^2 + \left(1 + \frac{M}{2r}\right)^4 (dx^2 + dy^2 + dz^2), \quad (3.3)$$

where $r = \sqrt{x^2 + y^2 + z^2}$.

3.3 Gauge conditions

Eq. (2.21)-(2.24) does not contain any time derivative of lapse function N nor of the shift vector β . This means that N and β are not dynamical variables. Therefore, we may choose the lapse and shift freely, without changing the physical solution g of the Einstein equation[6].

In this simulation, we choose the lapse function and the shift vector

$$N = \frac{1 - M/(2r)}{1 + M/(2r)}, \quad \beta = 0. \quad (3.4)$$

3.4 Initial data

After 3+1 decomposition, we should evolve forward in time some initial data. Instead of solving Hamiltonian and momentum constraint, we use well-known initial data from isotropic coordinates of Schwarzschild metric. So the initial spatial metric becomes $\gamma_{ij} = (1 + M/(2r))^4 \delta_{ij}$ and the initial extrinsic curvature becomes $K_{ij} = 0$ since there is no time-dependent and zero shift.

We will show that these data satisfies Eq. (2.21, 2.22). Since E and p_i are all 0 in vacuum space and $K_{ij} = 0$, the momentum constraint is naturally satisfied. Now, to satisfy the hamiltonian constraint, we need to show that the 3-metric Ricci scalar R is 0. Let's use a spherical coordinate system here. The non-vanishing Ricci tensor is:

$$R_{rr} = -\frac{8M}{r(M+2r)^2}, \quad (3.5)$$

$$R_{\theta\theta} = \frac{4Mr}{(M+2r)^2}, \quad (3.6)$$

$$R_{\phi\phi} = \frac{4Mr \sin^2(\theta)}{(M+2r)^2}. \quad (3.7)$$

Therefore, it can be seen that the 3-metric Ricci scalar $R = \gamma^{ij} R_{ij} = 0$, and it can be confirmed that the given constraint condition is well satisfied.

3.5 Numerical methods

3.5.1 Finite difference method

This simulation uses the finite difference method. The Taylor expansion of the function $f(x)$ in x_0 is

$$f(x_0 + h) = f(x_0) + \frac{f'(x_0)}{1!}h + \frac{f^{(2)}(x_0)}{2!}h^2 + \dots + \frac{f^{(n)}(x_0)}{n!}h^n + \dots. \quad (3.8)$$

Arranging this, we get

$$\begin{aligned}\frac{f(x_0 + h) - f(x_0)}{h} &= f'(x_0) + \frac{f^{(2)}(x_0)}{2!}h + \dots \\ &= f'(x_0) + \mathcal{O}(h).\end{aligned}\tag{3.9}$$

Accuracy is on the order of $\mathcal{O}(h)$.

The central difference method selects the function value from $x - h$ and $x + h$, respectively, and has a more accurate error of $\mathcal{O}(h^2)$.

$$f'(x_0) = \frac{f(x + h) - f(x - h)}{2h} + \mathcal{O}(h^2).\tag{3.10}$$

3.5.2 Boundary condition

The isotropic coordinates of the schwarzschild metric only describe the region outside the black hole horizon, i.e. $r = \frac{M}{2}$. Therefore, γ_{ij} and K_{ij} did not evolve in the region of $r \leq \frac{M}{2}$.

At the boundary outside the grid, the physical quantity at the corresponding point was calculated using linear extrapolation. For example, when we need to find the derivative at f_i ,

$$(\partial f)_{i-2} = \frac{f_{i-1} - f_{i-3}}{2h},\tag{3.11}$$

$$(\partial f)_{i-1} = \frac{f_i - f_{i-2}}{2h}.\tag{3.12}$$

Therefore, we get

$$\begin{aligned}(\partial f)_i &= \frac{f_{i-1} - f_{i-3}}{2h} \\ &= \frac{2f_i - f_{i-1} - 2f_{i-2} + f_{i-3}}{2h}.\end{aligned}\tag{3.13}$$

Alternatively, there is a way to use a fixed value at the grid boundary, as well as at the black hole horizon.

3.5.3 Inverse matrix

Cofactors were used to find the inverse matrix. See Appendix A.

3.6 Grid setting

The size of the grid is 100^3 , the mass of the black hole is $0.2M$, and the grid distance is $0.01M$. Thus, the horizon of a black hole corresponds to $0.1M$, i.e. the surface of a sphere with a radius of 10 grids.

3.7 Result

The results obtained at first do not change with time as shown in Fig. 3.1.

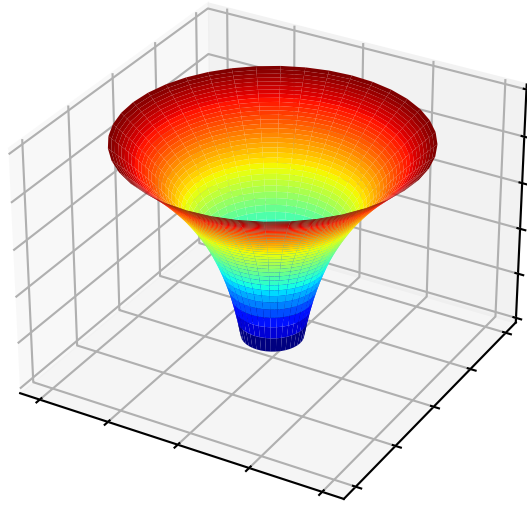


FIGURE 3.1: Flamm's paraboloid from the first simulation. It represents γ_{xx} in the equatorial plane where $\theta = \frac{\pi}{2}$. The slope at each point represents the magnitude of γ_{xx} .

However, when the errors in the coordinate system of the lapse function and minor errors were corrected, the metric could no longer be considered static as shown in Figure 1.

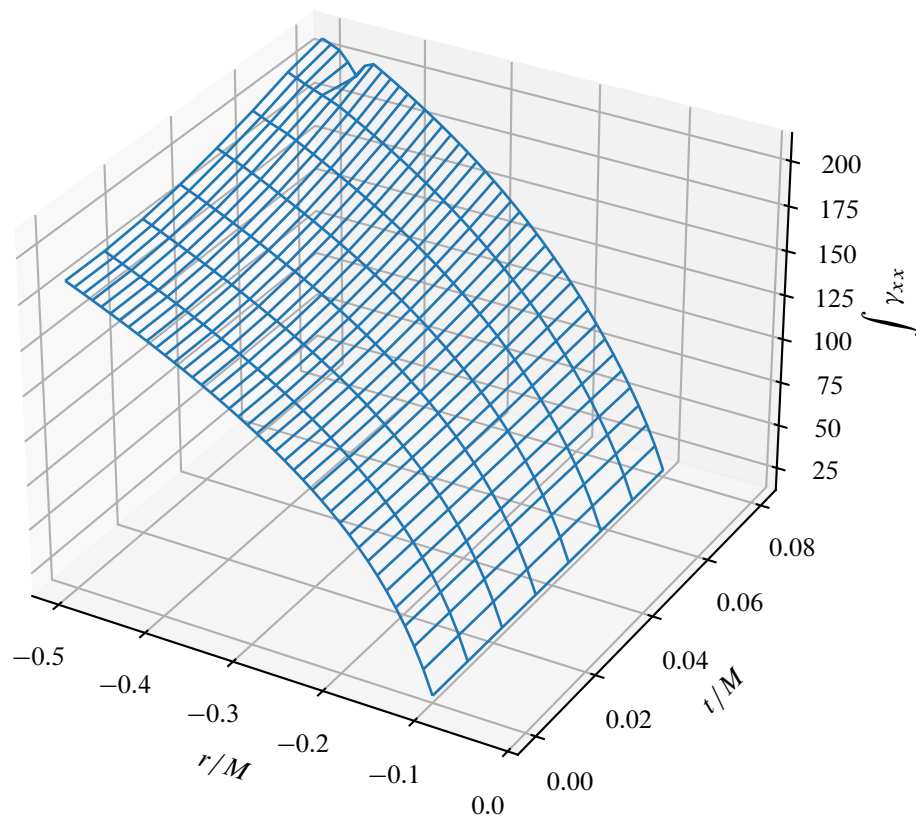


FIGURE 3.2: The unidirectional component of Flamm's paraboloid as a function of time. Even in a very short time interval, errors gradually accumulate over time, especially distortions can be seen.

Chapter 4

Discussion and Conclusion

4.1 Analytical calculation of evolution values

In order for the metric to be static according to time, $K_{ij} = 0$ must always be present. Also, to keep $K_{ij} = 0$, in Eq. (2.24), $-D_i D_j N + N R_{ij} = 0$. Let's check if this is true by calculating K_{rr} in the spherical coordinate system.

The lapse function is

$$N = \frac{1 - M/(2r)}{1 + M/(2r)} \quad (4.1)$$

And the covariant derivative of lapse function is

$$D_i D_j N = \frac{\partial^2 N}{\partial x^i \partial x^j} - \Gamma_{ij}^k \frac{\partial N}{\partial x^k}. \quad (4.2)$$

When $i = j = r$,

$$\frac{\partial^2 N}{\partial x^r \partial x^r} = -\frac{16M}{(M + 2r)^3}. \quad (4.3)$$

Since N is a function of r only, the significant term in the Christoffel symbols is

$$\Gamma_{rr}^r = -\frac{8M}{r(M + 2r)^2}. \quad (4.4)$$

Therefore, we get

$$D_r D_r N = -\frac{16M(Mr + 2M + 2r^2)}{r(M + 2r)^4}. \quad (4.5)$$

On the other hand, the Ricci tensor is

$$R_{rr} = -\frac{8M}{r(M + 2r)^2}, \quad (4.6)$$

we get

$$NR_{rr} = -\frac{8M(-M+2r)}{r(M+2r)^3}. \quad (4.7)$$

Therefore,

$$-D_r D_r N + NR_{rr} = \frac{8M^2(M+2r+4)}{r(M+2r)^4} \quad (4.8)$$

Contrary to expectations, it does not become 0.

4.2 How about initial data?

Let's check that the simulation calculates the Hamiltonian constraints well, at least for the initial data. In the case of the momentum constraint, since $K_{ij} = 0$, the condition is obviously satisfied.

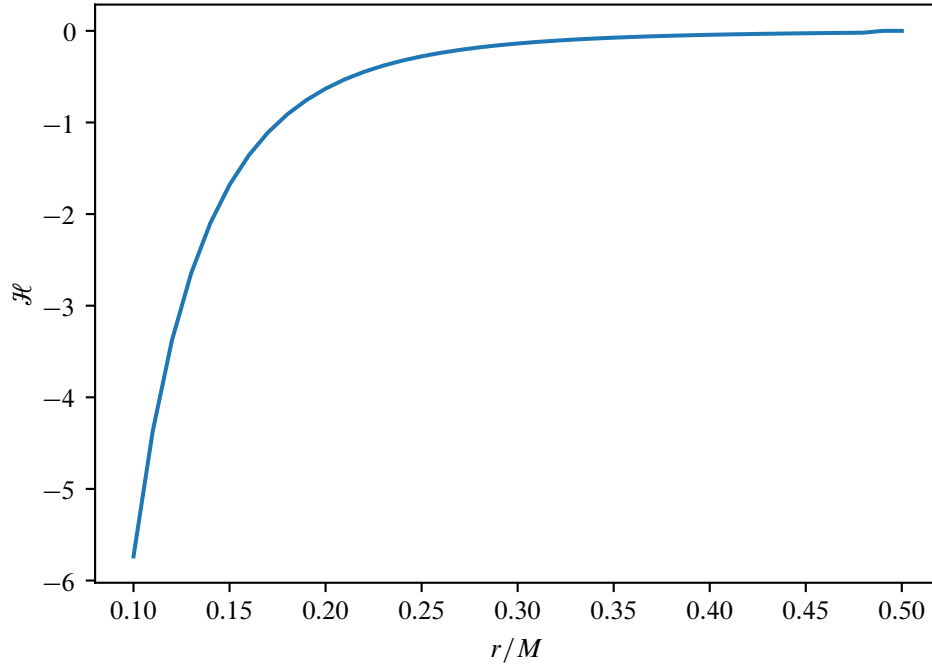


FIGURE 4.1: It shows the value of $\mathcal{H} \equiv R + K^2 - K_{ij}K^{ij}$, the left side of the Hamiltonian constraint, according to r .

It is close to the expected value of 0, but it can be seen that the error gets worse as you get closer to the black hole horizon. In this simulation, since the numerical derivative is calculated on the order of $\mathcal{O}(h^2)$, if the grid spacing is reduced from 0.01 to 0.005, that is, by a factor of 2, the error should be reduced by a factor of 4, which is shown in Figure 4.2.

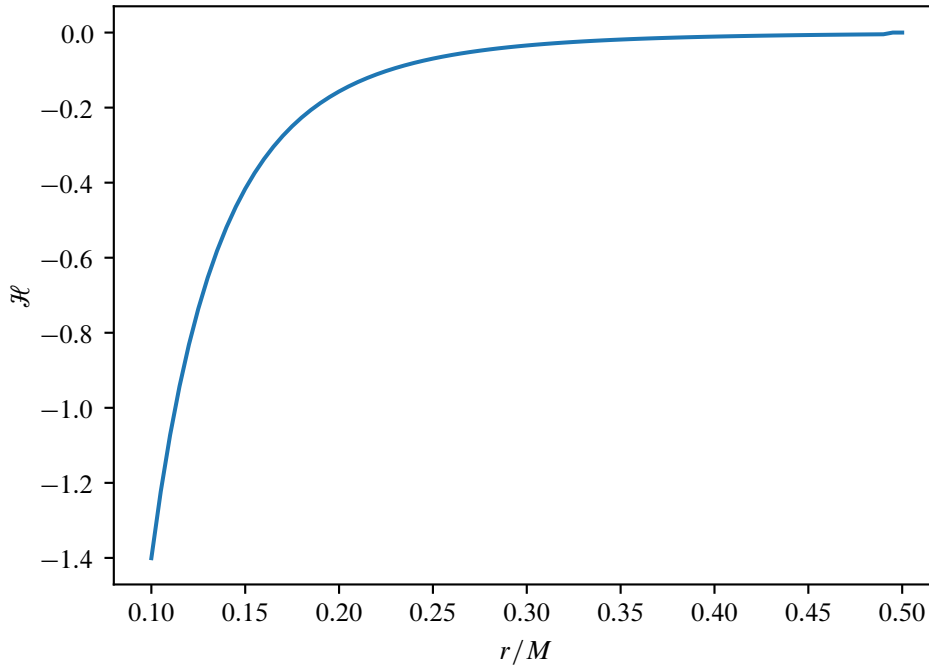


FIGURE 4.2: It shows the value of $\mathcal{H} \equiv R + K^2 - K_{ij}K^{ij}$, the left side of the Hamiltonian constraint, according to r . Compared to Fig. 4.1, the resolution has doubled.

4.3 Conclusion

The off-diagonal terms of γ_{ij} have a slight relative error of the order of 10^{-7} in comparison with diagonal terms. This accumulates over time. Ignoring these small errors, we were able to confirm that the spatial metric was static and still spherically symmetric. We should either accept a variation form such as the BSSN equation, or try to do more accurate numerical calculations for better simulations.

Appendix A

Frequently Asked Questions

A.1 Used code

This is part of the code used for simulation.

```
1  import pickle
2
3  SIZE = 101
4  CENTER = SIZE // 2
5  dx = 0.01
6  dt = 0.1
7  M = 0.2
8
9  to_ij = [[0, 0], [0, 1], [0, 2], [1, 1], [1, 2], [2, 2]]
10 _dx = [[1, 0, 0], [0, 1, 0], [0, 0, 1]]
11
12 def sqrt(x):
13     return x ** .5
14
15 def to_ind(x, y):
16     return to_ij.index(sorted([x, y]))
17
18
19 def dist(x, y, z):
20     return (x * x + y * y + z * z) ** .5
21
22
```

```

23 def psi(r):
24     return 1 + M / 2 / r
25
26
27 def lapse(r):
28     return (2 * r - M) / (2 * r + M)
29
30
31 def dNx(ind, x, y, z):
32     if ind == 1:
33         x, y = y, x
34     elif ind == 2:
35         x, z = z, x
36     return 4 * M * x / ((M + 2 * sqrt(x ** 2 + y ** 2 + z ** 2)) ** 2 *
    ↪ sqrt(x ** 2 + y ** 2 + z ** 2))
37 def dNxx(ind, x, y, z):
38     if ind == 3:
39         x, y = y, x
40     elif ind == 5:
41         x, z = z, x
42     return 4 * M * (-x ** 2 * (M + 2 * sqrt(x ** 2 + y ** 2 + z ** 2)) * (x
    ↪ ** 2 + y ** 2 + z ** 2) ** (3 / 2) - 4 * x ** 2 * (
43 x ** 2 + y ** 2 + z ** 2) ** 2 + (M + 2 * sqrt(x ** 2 + y ** 2 + z **
    ↪ 2)) * (
44 x ** 2 + y ** 2 + z ** 2) ** (5 / 2)) / (
45 (M + 2 * sqrt(x ** 2 + y ** 2 + z ** 2)) ** 3 * (x ** 2 + y ** 2 + z **
    ↪ 2) ** 3)
46 def dNxy(ind, x, y, z):
47     if ind == 2:
48         y, z = z, y
49     elif ind == 4:
50         x, z = z, x
51     return 4 * M * x * y * (-(M + 2 * sqrt(x ** 2 + y ** 2 + z ** 2)) * (x
    ↪ ** 2 + y ** 2 + z ** 2) - 4 * (

```

```

52 x ** 2 + y ** 2 + z ** 2) ** (3 / 2)) / (
53 (M + 2 * sqrt(x ** 2 + y ** 2 + z ** 2)) ** 3 * (x ** 2 + y ** 2 + z **
    ↪ 2) ** (5 / 2))
54 def ddlapse(ind, r, i, j, k):
55     _i, _j = to_ij[ind]
56     x, y, z = (i - CENTER) * dx, (j - CENTER) * dx, (k - CENTER) * dx
57     if _i == _j:
58         return dNxx(ind, x, y, z) - sum(Christoffel[_k][ind][i][j][k] * dNx(_k,
    ↪ x, y, z) for _k in range(3))
59     else:
60         return dNxy(ind, x, y, z) - sum(Christoffel[_k][ind][i][j][k] * dNx(_k,
    ↪ x, y, z) for _k in range(3))
61     # if ind == 0:
62     #     return -16 * M / (2 * r + M) ** 3 - 4 * M / (2 * r + M) ** 2 *
    ↪ Christoffel[0][0][i][j][k]
63
64
65 def transposeMatrix(m):
66     return list(map(list, zip(*m)))
67
68
69 def getMatrixMinor(m, i, j):
70     return [row[:j] + row[j + 1:] for row in (m[:i] + m[i + 1:])]
71
72
73 def getMatrixDeterminant(m):
74     if len(m) == 2:
75         return m[0][0] * m[1][1] - m[0][1] * m[1][0]
76     determinant = 0
77     for c in range(len(m)):
78         determinant += ((-1) ** c) * m[0][c] *
    ↪ getMatrixDeterminant(getMatrixMinor(m, 0, c))
79     return determinant
80

```

```

81
82 def getMatrixInverse(m):
83     determinant = getMatrixDeterminant(m)
84     if determinant == 0:
85         print('Warning : determinant is zero.')
86         return [[0] * 3 for _ in range(3)]
87     if len(m) == 2:
88         return [[m[1][1] / determinant, -1 * m[0][1] / determinant],
89                 [-1 * m[1][0] / determinant, m[0][0] / determinant]]
90     cofactors = []
91     for r in range(len(m)):
92         cofactorRow = []
93         for c in range(len(m)):
94             minor = getMatrixMinor(m, r, c)
95             cofactorRow.append(((−1) ** (r + c)) * getMatrixDeterminant(minor))
96         cofactors.append(cofactorRow)
97     cofactors = transposeMatrix(cofactors)
98     for r in range(len(cofactors)):
99         for c in range(len(cofactors)):
100             cofactors[r][c] = cofactors[r][c] / determinant
101     return cofactors
102
103 initial = 0
104 tc = 0
105 errest = 1
106 evolve = 0
107 for _tt in range(1):
108     print('Step :', tc)
109     # Orr, 1rtheta, 2rphi, 3thetatheta, 4thetaphi, 5phiphi
110
111     if initial:
112         gamma = [[[[0] * SIZE for _i in range(SIZE)] for _j in range(SIZE)] for
113                  ↪ _k in range(6)]

```

```

113 K = [[[[0] * SIZE for _i in range(SIZE)] for _j in range(SIZE)] for _k
    ↪ in range(6)]
114 else:
115 with open("gamma_%d.txt" % tc, "rb") as f:
116 gamma = pickle.load(f)
117 with open("K_%d.txt" % tc, "rb") as f:
118 K = pickle.load(f)
119
120 gamma_inv = [[[[0] * SIZE for _i in range(SIZE)] for _j in range(SIZE)]
    ↪ for _k in range(6)]
121 pgamma = [[[[0] * SIZE for _i in range(SIZE)] for _j in range(SIZE)]
    ↪ for _k in range(6)] for _ in range(3)]
122 meanK = [[0] * SIZE for _i in range(SIZE)] for _j in range(SIZE)]
123 R = [[[[0] * SIZE for _i in range(SIZE)] for _j in range(SIZE)] for _k
    ↪ in range(6)]
124
125 r0 = M / 2
126 if initial:
127 print('Set Inital Gamma')
128 for i in range(SIZE):
129 for j in range(SIZE):
130 for k in range(SIZE):
131 if i == j == k == CENTER:
132 continue
133 else:
134 r = dist(i - CENTER, j - CENTER, k - CENTER) * dx
135 gamma[0][i][j][k] = psi(r) ** 4
136 gamma[3][i][j][k] = psi(r) ** 4
137 gamma[5][i][j][k] = psi(r) ** 4
138 with open('gamma_%d.txt' % (0), 'wb') as f:
139 pickle.dump(gamma, f)
140 with open('K_%d.txt' % (0), 'wb') as f:
141 pickle.dump(K, f)
142

```

```

143 print('Get Inverse Gamma')
144 for i in range(SIZE):
145     for j in range(SIZE):
146         for k in range(SIZE):
147             if i == j == k == CENTER:
148                 continue
149
150     _subgamma = [[gamma[0][i][j][k], gamma[1][i][j][k], gamma[2][i][j][k]],
151                 [gamma[1][i][j][k], gamma[3][i][j][k], gamma[4][i][j][k]],
152                 [gamma[2][i][j][k], gamma[4][i][j][k], gamma[5][i][j][k]]]
153     _gammainv = getMatrixInverse(_subgamma)
154     for _ind in range(6):
155         gamma_inv[_ind][i][j][k] = _gammainv[to_ij[_ind][0]][to_ij[_ind][1]]
156
157 print('Get Mean K')
158 for i in range(SIZE):
159     for j in range(SIZE):
160         for k in range(SIZE):
161             if i == j == k == CENTER:
162                 continue
163             meanK[i][j][k] = sum(
164                 gamma_inv[to_ind(_i, _j)][i][j][k] * K[to_ind(_i, _j)][i][j][k] for _i
165                 ↪ in range(3) for _j in range(3))
166
167 print('Get Partial Gamma')
168 for i in range(SIZE):
169     for j in range(SIZE):
170         for k in range(SIZE):
171             if abs(i - CENTER) + abs(j - CENTER) + abs(k - CENTER) <= 1:
172                 continue
173             if i == 0 or i == SIZE - 1 or j == 0 or j == SIZE - 1 or k == 0 or k ==
174                 ↪ SIZE - 1:
175                 continue

```

```

175 for _sup in range(6):
176     # _ipgamma = [(gamma[_sup][i + 1][j][k] - gamma[_sup][i - 1][j][k]) /
    ↪ dx / 2,
177     #             (gamma[_sup][i][j + 1][k] - gamma[_sup][i][j - 1][k]) /
    ↪ dx / 2,
178     #             (gamma[_sup][i][j][k + 1] - gamma[_sup][i][j][k - 1]) /
    ↪ dx / 2]
179 for _sub in range(3):
180     # pgamma[_sub][_sup][i][j][k] = _ipgamma[_sub]
181     # _case = 0
182     # _coord = [i, j, k]
183     # _fixed = _coord[_sub]
184     # _coordp = _coord[:]
185     # _coordp[_sub] += 1
186     # _coordm = _coord[:]
187     # _coordm[_sub] -= 1
188     # if _fixed == SIZE - 1 or _coordp == [CENTER, CENTER, CENTER]:
189     #     _case = 2
190     # elif _fixed == 0 or _coordm == [CENTER, CENTER, CENTER]:
191     #     _case = 1
192     pgamma[_sub][_sup][i][j][k] = (gamma[_sup][i + _dx[_sub][0]][j +
    ↪ _dx[_sub][1]][k + _dx[_sub][2]]
193     - gamma[_sup][i - _dx[_sub][0]][j - _dx[_sub][1]][k - _dx[_sub][2]]) /
    ↪ dx / 2
194     # elif _case == 1:
195     #     pgamma[_sub][_sup][i][j][k] = (2 * gamma[_sup][i + 2 *
    ↪ _dx[_sub][0]][j + 2 * _dx[_sub][1]][k + 2 * _dx[_sub][2]]
196     #     - gamma[_sup][i + 3 *
    ↪ _dx[_sub][0]][j + 3 * _dx[_sub][1]][k + 3 * _dx[_sub][2]]
197     #     - 2 * gamma[_sup][i][j][k]
198     #     + gamma[_sup][i +
    ↪ _dx[_sub][0]][j + _dx[_sub][1]][k + _dx[_sub][2]]) / dx / 2
199     # else:

```

```

200 #      pgamma[_sub][_sup][i][j][k] = (- 2 * gamma[_sup][i - 2 *
    ↳ _dx[_sub][0]][j - 2 * _dx[_sub][1]][
201 #                                     k - 2 * _dx[_sub][2]]
202 #                                     + gamma[_sup][i - 3 *
    ↳ _dx[_sub][0]][j - 3 * _dx[_sub][1]][
203 #                                     k - 3 * _dx[_sub][2]]
204 #                                     + 2 * gamma[_sup][i][j][k]
205 #                                     - gamma[_sup][i -
    ↳ _dx[_sub][0]][j - _dx[_sub][1]][
206 #                                     k - _dx[_sub][2]]) / dx /
    ↳ 2
207
208 if initial:
209     with open('Pgamma.txt', 'wb') as f:
210         pickle.dump(pgamma, f)
211
212     print('Get Christoffel')
213     Christoffel = [[[[[0] * SIZE for _i in range(SIZE)] for _j in
    ↳ range(SIZE)] for _k in range(6)] for _l in range(3)]
214     for i in range(SIZE):
215         print('Christoffel for : ', i)
216         for j in range(SIZE):
217             for k in range(SIZE):
218                 if abs(i - CENTER) + abs(j - CENTER) + abs(k - CENTER) <= 1:
219                     continue
220                 if i == 0 or i == SIZE - 1 or j == 0 or j == SIZE - 1 or k == 0 or k ==
    ↳ SIZE - 1:
221                     continue
222
223                 for _sup in range(3):
224                     for _sub in range(6):
225                         # _subgamma = [[gamma[0][i][j][k], gamma[1][i][j][k],
    ↳ gamma[2][i][j][k]],

```



```

226 #          [gamma[1][i][j][k], gamma[3][i][j][k],
    ↪ gamma[4][i][j][k]],
227 #          [gamma[2][i][j][k], gamma[4][i][j][k],
    ↪ gamma[5][i][j][k]]]
228 # _gammainv = getMatrixInverse(_subgamma)
229 _subi, _subj = to_ij[_sub]
230 Christoffel[_sup][_sub][i][j][k] = sum(
231 gamma_inv[to_ind(1, _sup)][i][j][k] * (pgamma[_subi][to_ind(1,
    ↪ _subj)][i][j][k] +
232 pgamma[_subj][to_ind(_subi, 1)][i][j][k] -
233 pgamma[1][to_ind(_subi, _subj)][i][j][k]) for l in
234 range(3)) / 2
235 if initial:
236 with open('Christoffel.txt', 'wb') as f:
237 pickle.dump(Christoffel, f)
238
239 print('Get Ricci Tensor')
240
241 def pChristoffel(sup, sub, i, j, k, x):
242 # _case = 0
243 # _coord = [i, j, k]
244 # _fixed = _coord[x]
245 # _coordp = _coord[:]
246 # _coordp[x] += 1
247 # _coordm = _coord[:]
248 # _coordm[x] -= 1
249 # if _fixed == SIZE - 1 or _coordp == [CENTER, CENTER, CENTER]:
250 #     _case = 2
251 # elif _fixed == 0 or _coordm == [CENTER, CENTER, CENTER]:
252 #     _case = 1
253
254 # if _case == 0:

```

```

255 return (Christoffel[sup][sub][i + _dx[x][0]][j + _dx[x][1]][k +
    ↪ _dx[x][2]] - Christoffel[sup][sub][i - _dx[x][0]][j - _dx[x][1]][k
    ↪ - _dx[x][2]]) / dx / 2
256 # elif _case == 1:
257 #     return (2 * Christoffel[sup][sub][i + 2 * _dx[x][0]][j + 2 *
    ↪ _dx[x][1]][k + 2 * _dx[x][2]]
258 #         - Christoffel[sup][sub][i + 3 * _dx[x][0]][j + 3 *
    ↪ _dx[x][1]][k + 3 * _dx[x][2]]
259 #         - 2 * Christoffel[sup][sub][i][j][k]
260 #         + Christoffel[sup][sub][i + _dx[x][0]][j + _dx[x][1]][k +
    ↪ _dx[x][2]])
261 # else:
262 #     return (- 2 * Christoffel[sup][sub][i - 2 * _dx[x][0]][j - 2 *
    ↪ _dx[x][1]][k - 2 * _dx[x][2]]
263 #         + Christoffel[sup][sub][i - 3 * _dx[x][0]][j - 3 *
    ↪ _dx[x][1]][k - 3 * _dx[x][2]]
264 #         + 2 * Christoffel[sup][sub][i][j][k]
265 #         - Christoffel[sup][sub][i - _dx[x][0]][j - _dx[x][1]][k -
    ↪ _dx[x][2]])
266
267
268
269 for i in range(SIZE):
270     print('Ricci Tensor for :', i)
271     for j in range(SIZE):
272         for k in range(SIZE):
273             if abs(i - CENTER) + abs(j - CENTER) + abs(k - CENTER) <= 2:
274                 continue
275             # if i == j == k == CENTER or i == 0 or i == SIZE - 1 or j == 0 or j ==
    ↪ SIZE - 1 or k == 0 or k == SIZE - 1:
276                 # continue
277             if i <= 1 or i >= SIZE - 2 or j <= 1 or j >= SIZE - 2 or k <= 1 or k >=
    ↪ SIZE - 2:
278                 continue

```

```

279
280 for _ind in range(6):
281     _i, _j = to_ij[_ind]
282     R[_ind][i][j][k] = sum(pChristoffel(_k, _ind, i, j, k, _k) for _k in
        ↪ range(3)) \
283     - sum(pChristoffel(_k, to_ind(_i, _k), i, j, k, _j) for _k in
        ↪ range(3)) \
284     + sum(
285     Christoffel[_k][to_ind(_i, _j)][i][j][k] * Christoffel[_l][to_ind(_k,
        ↪ _l)][i][j][k] for _k in
286     range(3) for _l in range(3)) \
287     - sum(
288     Christoffel[_l][to_ind(_i, _k)][i][j][k] * Christoffel[_k][to_ind(_l,
        ↪ _j)][i][j][k] for _k in
289     range(3) for _l in range(3))
290
291 if errest:
292     print('Error est')
293     print('mean R')
294     meanR = [[0] * SIZE for _i in range(SIZE)] for _j in range(SIZE)]
295     for i in range(SIZE):
296         for j in range(SIZE):
297             for k in range(SIZE):
298                 if i == j == k == CENTER:
299                     continue
300                 meanR[i][j][k] = sum(
301                 gamma_inv[to_ind(_i, _j)][i][j][k] * R[to_ind(_i, _j)][i][j][k] for _i
        ↪ in range(3) for _j in range(3))
302     print('KK')
303     KK = [[0] * SIZE for _i in range(SIZE)] for _j in range(SIZE)]
304     for i in range(SIZE):
305         for j in range(SIZE):
306             for k in range(SIZE):
307                 if i == j == k == CENTER:

```

```

308 continue
309 KK[i][j][k] = sum(sum(gamma_inv[to_ind(_i, _mu)][i][j][k] *
    ↪ gamma_inv[to_ind(_j, _nu)][i][j][k] *
310 K[to_ind(_mu, _nu)][i][j][k] for _mu in range(3) for _nu in range(3))
311 * K[to_ind(_i, _j)][i][j][k] for _i in range(3) for _j in range(3))
312 err = [[[0] * SIZE for _i in range(SIZE)] for _j in range(SIZE)]
313 for i in range(SIZE):
314 for j in range(SIZE):
315 for k in range(SIZE):
316 if i == j == k == CENTER:
317 continue
318 err[i][j][k] = meanR[i][j][k] + (meanK[i][j][k]) ** 2 - KK[i][j][k]
319 with open('error_%d.txt' % (tc), 'wb') as f:
320 pickle.dump(err, f)
321
322
323
324
325 # R[_ind][i][j][k] = sum((Christoffel[_k][_ind][i + _dx[_k][0]][j +
    ↪ _dx[_k][1]][k + _dx[_k][2]] -
326 #                               Christoffel[_k][_ind][i][j][k]) / dx for _k
    ↪ in range(3)) \
327 #                               - sum((Christoffel[_k][to_ind(_i, _k)][i +
    ↪ _dx[_j][0]][j + _dx[_j][1]][
328 #                               k + _dx[_j][2]] -
    ↪ Christoffel[_k][to_ind(_i, _k)][i][j][k]) / dx for _k in
329 #                               range(3)) \
330 #                               + sum(
331 #                               Christoffel[_k][to_ind(_i, _j)][i][j][k] *
    ↪ Christoffel[_l][to_ind(_k, _l)][i][j][k] for _k in
332 #                               range(3) for _l in range(3)) \
333 #                               - sum(
334 #                               Christoffel[_l][to_ind(_i, _k)][i][j][k] *
    ↪ Christoffel[_k][to_ind(_l, _j)][i][j][k] for _k in

```

```

335 #     range(3) for _l in range(3))
336
337 # with open("gamma.txt", 'wb') as f:
338 #     pickle.dump(gamma, f)
339 # with open('gamma_inv.txt', 'wb') as f:
340 #     pickle.dump(gamma_inv, f)
341 # with open('pgamma.txt', 'wb') as f:
342 #     pickle.dump(pgamma, f)
343 # with open('K.txt', 'wb') as f:
344 #     pickle.dump(K, f)
345 # with open('meanK.txt', 'wb') as f:
346 #     pickle.dump(meanK, f)
347 # with open('R.txt', 'wb') as f:
348 #     pickle.dump(R, f)
349 # with open('Christoffel.txt', 'wb') as f:
350 #     pickle.dump(Christoffel, f)
351 #
352 if evolve:
353     print('Evolve')
354     _gamma = [[[0] * SIZE for _i in range(SIZE)] for _j in range(SIZE)]
355     ↪ for _k in range(6)]
356     _K = [[[0] * SIZE for _i in range(SIZE)] for _j in range(SIZE)] for _k
357     ↪ in range(6)]
358
359 for i in range(SIZE):
360     print('Evolve for :', i)
361     for j in range(SIZE):
362         for k in range(SIZE):
363             # if i == j == k == CENTER:
364             #     continue
365             # if i == 0 or i == SIZE - 1 or j == 0 or j == SIZE - 1 or k == 0 or k
366             ↪ == SIZE - 1:
367             #     continue

```

```

365 if i <= 1 or i >= SIZE - 2 or j <= 1 or j >= SIZE - 2 or k <= 1 or k >=
    ↪ SIZE - 2:
366 for ind in range(6):
367     _gamma[ind][i][j][k] = gamma[ind][i][j][k]
368     _K[ind][i][j][k] = K[ind][i][j][k]
369     continue
370     r = dist(i - CENTER, j - CENTER, k - CENTER) * dx
371     if r <= r0:
372         for ind in range(6):
373             _gamma[ind][i][j][k] = gamma[ind][i][j][k]
374             _K[ind][i][j][k] = K[ind][i][j][k]
375             continue
376         for ind in range(6):
377             _i, _j = to_ij[ind]
378             _gamma[ind][i][j][k] = gamma[ind][i][j][k] - 2 * lapse(r) *
    ↪ K[ind][i][j][k] * dt
379             _K[ind][i][j][k] = K[ind][i][j][k] - dt * (-ddlapse(ind, r, i, j, k) +
    ↪ lapse(r) * (R[ind][i][j][k]
380 + meanK[i][j][k] *
381 K[ind][i][j][
382 k] - 2 * sum(
383 K[to_ind(_i, _k)][i][j][k] * sum(
384 gamma_inv[to_ind(_k, _u)][i][j][k] * K[to_ind(_u, _j)][i][j][k] for _u
    ↪ in range(3)) for
385 _k in range(3))))
386
387 with open('gamma_%d.txt' % (tc + 1), 'wb') as f:
388     pickle.dump(_gamma, f)
389     with open('K_%d.txt' % (tc + 1), 'wb') as f:
390         pickle.dump(_K, f)
391     tc += 1

```

Bibliography

- [1] Charles W Misner, Kip S Thorne, and John Archibald Wheeler. *Gravitation*. Macmillan, 1973.
- [2] Bernd Brügmann. “Adaptive mesh and geodesically sliced Schwarzschild space-time in 3+ 1 dimensions”. In: *Physical Review D* 54.12 (1996), p. 7361.
- [3] Miguel Alcubierre. *Introduction to 3+ 1 numerical relativity*. Vol. 140. OUP Oxford, 2008.
- [4] Thomas W Baumgarte and Stuart L Shapiro. *Numerical relativity: solving Einstein’s equations on the computer*. Cambridge University Press, 2010.
- [5] Robert M Wald. *General relativity*. University of Chicago press, 2010.
- [6] Eric Gourgoulhon. *3+ 1 formalism in general relativity: bases of numerical relativity*. Vol. 846. Springer Science & Business Media, 2012.
- [7] Anthony Zee. *Einstein gravity in a nutshell*. Vol. 14. Princeton University Press, 2013.
- [8] WMHP Vindana and KAIL Wijewardena Gamalath. “Simulation of rotating black holes”. In: *World Scientific News* 114 (2018), pp. 106–125.
- [9] Bernard Schutz. *A first course in general relativity*. Cambridge university press, 2022.