

# Cosmology Tutorial 9

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## 1

### 1.1

In the radiation era fluctuations are frozen into a high density radiation field.

Radiation	Plasma	Neutral
=	=	+
=	=	+
=	+	+

Table 1: Baryonic Isothermal

### 1.2

Silk damping occurs in the radiation era but Silk mass is reduced after decoupling.

Radiation	Plasma	Neutral
-	-	+
+	+	+
+	+	+

Table 2: Baryonic Adiabatic

### 1.3

Fluctuations on small scales remain undamped.

Radiation	Plasma	Neutral
+	+	+
-	+	+
-	+	+

Table 3: CDM

## 1.4

Damping occurs due to free-streaming when particles are relativistic before  $t_{eq}$ .

Radiation	Plasma	Neutral
-	+	+
-	+	+
-	+	+

Table 4: HDM

## 2

$$k_x = n_x \frac{2\pi}{100}$$

$$2 \frac{2\pi}{100} = .126$$

$$3 \frac{2\pi}{100} = .188$$

so  $n_x$  must be either 0,2 or 3. The wavelengths are:

$$\frac{2\pi}{2} = \pi$$

$$\frac{2\pi}{3} = 2.094$$

There are  $3^3 = 27$  permutations of  $n_x$ ,  $n_y$  and  $n_z$ . For  $1 < k < 2$  the interval is 10 times larger so presumably there are 10 times as many permutations i.e. 270? .

### 3

#### 3.1

The Horizon scale at matter-energy equality is approximately 100 Mpc. The amplitude of fluctuations is  $\approx 10^{-4}$ .

#### 3.2

Redshift	Amplitude	Log(Amplitude)	Log(a)
$4 \times 10^7$	$10^{-8}$	-8	-7.6
$6.6 \times 10^6$	$10^{-7}$	-7	-6.8
$2.2 \times 10^6$	$10^{-6}$	-6	-6.3
$1.4 \times 10^6$	$10^{-5}$	-5	-6.1
$6.3 \times 10^5$	$10^{-4}$	-4	-5.7
39000	$10^{-3}$	-3	-4.6
490	.01	-2	-2.7
29	.1	-1	-1.5
2.2	1	0	-0.5

Table 5: Question 3 data

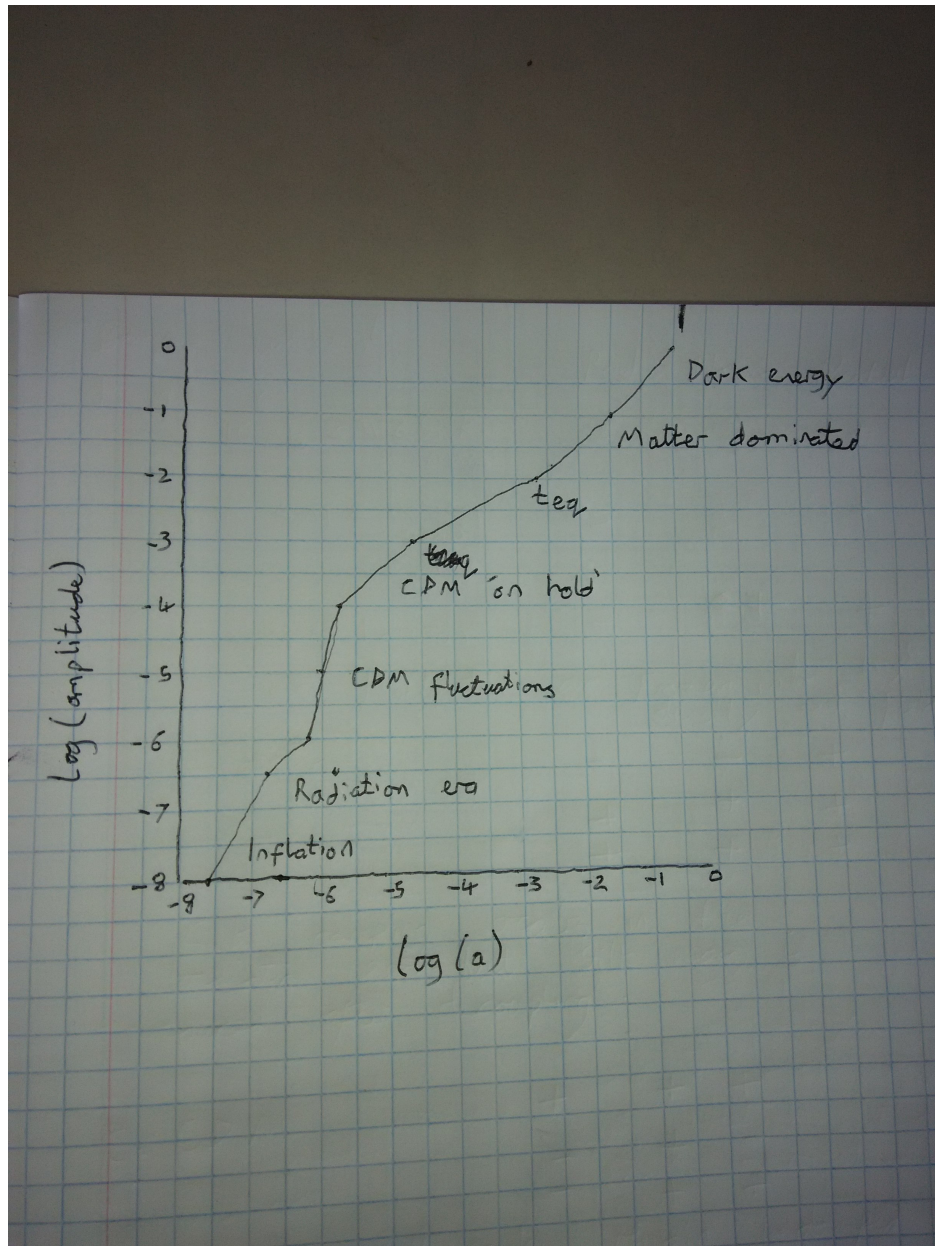


Figure 1:

4

Horizon problem - regions of the sky more than  $1^\circ (\approx 200 Kpc)$  were never in causal contact before decoupling so no information could have been exchanged. So why it is so isotropic?

Flatness problem - Why is  $\Omega = 1$  in universe today? The early universe must have had  $\Omega = 1$  as well (proof is in the notes).

Inflation solves the horizon problem because the causal horizon could have been larger in the past. Inflation also predicts that  $\Omega = 1$  because size of universe was so large after inflation that any curvature would have been negligible.

## 5

$$K_0 = 1 + .227N_\nu$$

$$N_\nu = 3.046$$

$$K_0 = 1.691$$

$$10^{-35} = 1.691^{-\frac{1}{2}} \left( \frac{1.52 \times 10^{10}}{T} \right)^2$$

$$\left( \frac{1.52 \times 10^{10}}{T} \right)^2 = \frac{10^{-35}}{.769}$$

$$\frac{1.52 \times 10^{10}}{T} = \sqrt{1.3 \times 10^{-35}}$$

$$T = \frac{1.52 \times 10^{10}}{3.6 \times 10^{-18}}$$

$$T = 4.22 \times 10^{27} K$$

$$kT = 8.62 \times 10^{-5} eV K^{-1} \times 4.22 \times 10^{27} K$$

$$= 3.64 \times 10^{23} eV$$

$$= 3.64 \times 10^{14} GeV$$

I couldn't answer the last bit.