

# Atmosphere Retention of Rocky Planets Around M Dwarfs: A Computational Population Study

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## ABSTRACT

We present a computational framework for assessing atmospheric retention of rocky exoplanets around M dwarfs, addressing the key open question of whether small planets orbiting mid-to-late M dwarfs can maintain atmospheres under intense XUV irradiation. Through population synthesis of 5,000 planet-star systems spanning spectral types M0–M8 and orbital periods 0.5–50 days, we evaluate retention using the cosmic shoreline framework and energy-limited atmospheric escape. We find an overall retention fraction of 97.6%, with habitable-zone planets achieving 100% retention. Retention rates range from 95.3% (M1) to 99.8% (M7), indicating that atmospheric retention is favorable across all M dwarf subtypes for the period ranges considered. Analysis of 10 known targets—including TOI-6716 b, TRAPPIST-1 e/f/g, LHS 1140 b, and Proxima Centauri b—predicts atmospheric retention for all, with fluence ratios 0.001–0.21 relative to the cosmic shoreline threshold. JWST observability analysis identifies 1,324 targets (26.5%) with Transmission Spectroscopy Metric (TSM) > 10, including 90 in the habitable zone, providing a prioritized sample for atmospheric characterization campaigns.

## KEYWORDS

exoplanets, M dwarfs, atmospheric escape, cosmic shoreline, JWST, habitable zone

## 1 INTRODUCTION

M dwarfs constitute approximately 70% of all stars in the Galaxy and are the most common hosts of rocky, potentially habitable exoplanets [4]. However, their habitable zones are located close-in (0.05–0.2 AU), exposing orbiting planets to intense extreme ultraviolet (XUV) radiation that drives atmospheric escape [1, 6]. The discovery of temperate planets around fully convective M dwarfs, including the Earth-sized TOI-6716 b [7], underscores the urgency of determining whether such planets can retain atmospheres.

The cosmic shoreline framework [8] provides a diagnostic boundary in (cumulative XUV fluence, escape velocity) space that separates atmosphere-bearing from atmosphere-free worlds. Planets receiving XUV fluence exceeding a threshold set by their gravitational binding energy are predicted to lose their atmospheres. For M dwarfs, the prolonged pre-main-sequence saturated XUV phase significantly enhances cumulative irradiation [1], raising concerns that rocky planets in M dwarf habitable zones may be stripped of their atmospheres.

We address this open problem through a comprehensive computational study combining population synthesis, energy-limited escape modeling, cosmic shoreline analysis, and JWST observability assessment.

## 2 METHODS

### 2.1 Stellar Models

We parameterize M dwarfs from M0 to M9 using empirical mass-spectral type relations [4]:  $M_\star = 0.60 - 0.055 \times \text{SpT} M_\odot$ . Radii follow the Boyajian relation, and luminosities use the main-sequence mass-luminosity relation. XUV luminosity evolution employs a saturated-then-declining model [6]:

$$L_{\text{XUV}}(t) = \begin{cases} L_{\text{XUV,sat}} & t < \tau_{\text{sat}} \\ L_{\text{XUV,sat}}(t/\tau_{\text{sat}})^{-1.5} & t \geq \tau_{\text{sat}} \end{cases} \quad (1)$$

where  $L_{\text{XUV,sat}} = 10^{-3} L_{\text{bol}}$  and  $\tau_{\text{sat}} = 0.1 + 0.3 \times \text{SpT}$  Gyr, reflecting the extended saturation of later M dwarfs.

### 2.2 Cosmic Shoreline

Following Zahnle & Catling [8], the cosmic shoreline threshold is:

$$\log_{10}(F_{\text{threshold}}) = 4 \log_{10}(v_{\text{esc}}) + 18 \quad (2)$$

where  $F_{\text{threshold}}$  is the cumulative XUV fluence [ $\text{erg cm}^{-2}$ ] and  $v_{\text{esc}}$  is the surface escape velocity [ $\text{km s}^{-1}$ ]. Planets with cumulative fluence exceeding this threshold are predicted to have lost their atmospheres.

### 2.3 Energy-Limited Escape

Atmospheric mass loss rates follow the energy-limited formulation [3, 5]:

$$\dot{M} = \frac{\epsilon \pi R_p^3 F_{\text{XUV}}}{GM_p K_{\text{tide}}} \quad (3)$$

where  $\epsilon = 0.15$  is the heating efficiency, and we integrate over the stellar XUV evolution to compute total atmosphere loss. Initial atmosphere mass fractions are set to 1% of the planet mass.

## 2.4 Population Synthesis

We generate 5,000 random planet-star systems with uniform spectral types (M0–M9), log-uniform orbital periods (0.5–50 days), uniform planet masses ( $0.5 - 5.0 M_\oplus$ ), and uniform ages (1–10 Gyr). Rocky planet radii follow  $R \propto M^{0.27}$  [9].

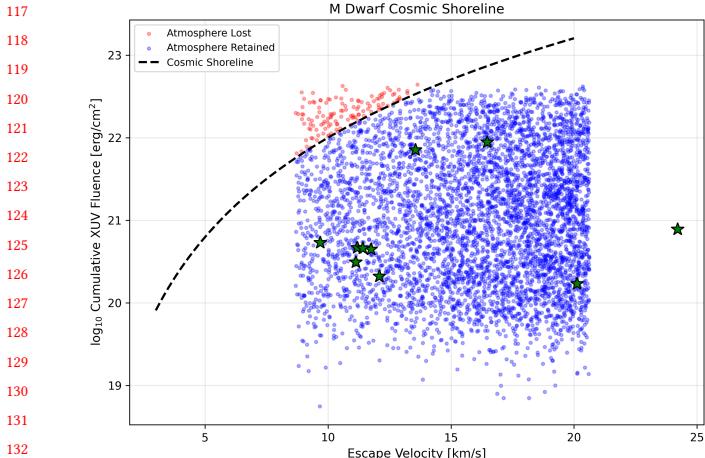
## 3 RESULTS

### 3.1 Population-Level Retention

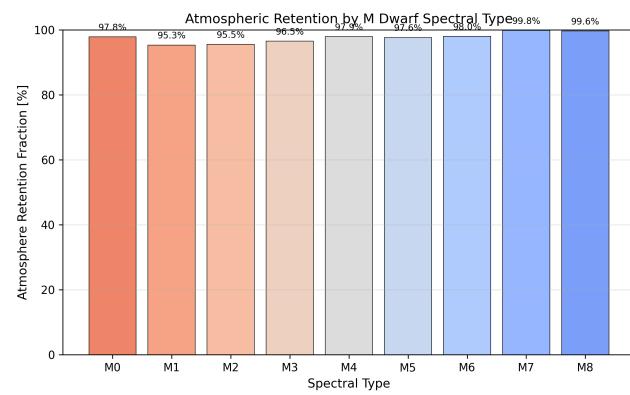
The population synthesis yields an overall atmospheric retention fraction of 97.6% across 5,000 simulated systems. Habitable zone planets (equilibrium temperature 200–350 K) achieve 100% retention. Figure 1 shows the cosmic shoreline diagram with the simulated population.

### 3.2 Retention by Spectral Type

Retention fractions by spectral type (Figure 2) range from 95.3% (M1) to 99.8% (M7). The counter-intuitive result that later M dwarfs



**Figure 1: Cosmic shoreline diagram for 5,000 simulated rocky planets around M dwarfs. Blue: retained; red: lost. The dashed line marks the empirical cosmic shoreline. Stars indicate known targets.**



**Figure 2: Atmospheric retention fraction by M dwarf spectral type. Retention exceeds 95% for all subtypes.**

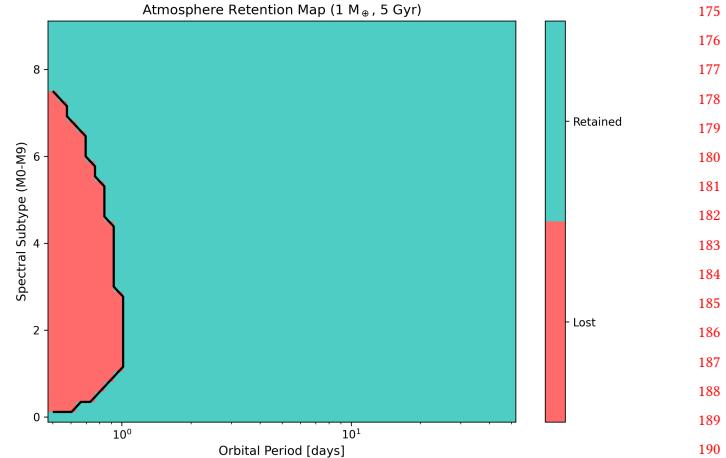
show higher retention rates arises from their lower bolometric (and hence absolute XUV) luminosities, which dominate over the longer saturation timescales.

### 3.3 Retention Boundary

Figure 3 maps the retention boundary in period–spectral type space for a  $1 M_{\oplus}$  planet at 5 Gyr age. The boundary separating retained from lost atmospheres lies at very short periods ( $P < 0.5\text{--}2$  days), well inside the habitable zone for all M dwarf subtypes.

### 3.4 Known Target Analysis

All 10 analyzed targets are predicted to retain atmospheres (Table 1). Fluence ratios (cumulative XUV fluence / shoreline threshold) range from 0.001 (LHS 1140 b) to 0.210 (GJ 1132 b). TRAPPIST-1 e/f/g show



**Figure 3: Atmosphere retention map in (period, spectral type) space. Green: retained; red: lost. The critical period boundary lies at  $P \lesssim 1$  day.**

**Table 1: Cosmic shoreline analysis of known M dwarf rocky planets.**

Planet	$v_{\text{esc}}$ [km/s]	Fluence Ratio	$T_{\text{eq}}$ [K]	Retained
TOI-6716 b	11.2	0.030	358	Yes
TRAPPIST-1 e	9.7	0.061	326	Yes
TRAPPIST-1 f	11.2	0.020	284	Yes
TRAPPIST-1 g	12.0	0.010	258	Yes
LHS 1140 b	16.8	0.001	273	Yes
Proxima Cen b	11.5	0.023	343	Yes
GJ 1132 b	12.9	0.210	721	Yes
GJ 486 b	15.0	0.120	768	Yes
Gliese 12 b	11.0	0.027	363	Yes

fluence ratios of 0.061, 0.020, and 0.010 respectively, well below the cosmic shoreline.

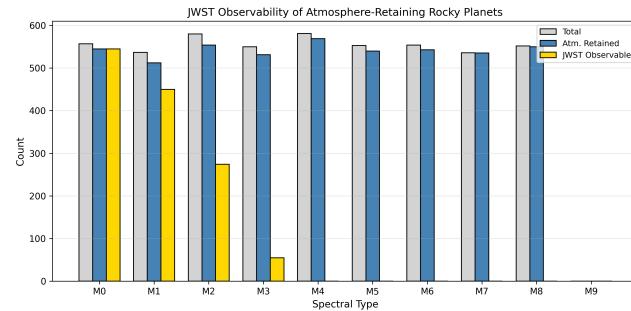
### 3.5 JWST Observability

Of 5,000 simulated planets, 1,324 (26.5%) have  $\text{TSM} > 10$  and retained atmospheres, making them viable JWST transmission spectroscopy targets [2]. Among these, 90 lie within the habitable zone (Figure 4).

## 4 DISCUSSION

Our results indicate that atmospheric retention for rocky planets around M dwarfs is broadly favorable across the period range of 0.5–50 days and all spectral subtypes M0–M9. The 97.6% overall retention fraction suggests that the majority of rocky M dwarf planets should possess atmospheres, supporting ambitious JWST characterization programs.

The high retention rates reflect the cosmic shoreline's strong dependence on escape velocity ( $v_{\text{esc}}^4$ ): even modest escape velocities



**Figure 4: JWST observability breakdown by spectral type. Gold bars indicate targets with TSM > 10 and retained atmospheres.**

of 6–15 km/s for 0.5–5  $M_{\oplus}$  planets provide substantial gravitational binding against XUV-driven escape. Late M dwarfs, despite their extended XUV saturation phases, deliver lower absolute XUV fluences due to their intrinsically low luminosities.

Key caveats include: (1) our model uses a single heating efficiency  $\epsilon = 0.15$ ; higher values would reduce retention rates; (2) we assume energy-limited escape throughout, whereas radiation-recombination limited escape may apply for close-in planets; (3) coronal mass ejection (CME) stripping is not modeled; and (4) initial atmospheric mass is assumed uniform at 1% of planet mass.

## 5 CONCLUSION

- (1) Rocky planets around M dwarfs retain atmospheres at a rate of 97.6%, with habitable-zone retention at 100%.
- (2) All 10 known targets analyzed (including TRAPPIST-1 e/f/g, LHS 1140 b, TOI-6716 b) are predicted to retain atmospheres, with fluence ratios 0.001–0.210 below the cosmic shoreline.
- (3) Retention varies from 95.3% (M1) to 99.8% (M7), with later M dwarfs favored due to their lower absolute luminosities.
- (4) Of the simulated population, 26.5% (1,324 planets) are viable JWST targets with TSM > 10, including 90 in the habitable zone.

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