

Non-Dimensionalization of an Anelastic Stable–Unstable Layer in **Rayleigh**

Loren Matilsky

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1 Purpose

This document describes how to implement, using the **Rayleigh** code, a coupled stable–unstable system. More precisely, the system consists of an unstable convection zone (CZ) that lies adjacent to a stable radiative zone (RZ) in a spherical shell. If the CZ is on top of the RZ we describe the system as “solar-like” or a “tachocline,” etc. If the CZ is underneath the RZ (now considered a “weather layer”) we describe the system as “Jovian” or “Jupiter,” etc. It essentially the same problem (at least to specify) either way, making this document relevant to Nic & COFFIES (tachocline), and to Geoff, Nic, and Matt (Jupiter). In the following sections, we describe the nondimensionalization, governing equations, and reference state.

2 General Equations Solved in **Rayleigh**

In general (with rotation and magnetism), **Rayleigh** time-evolves a set of coupled PDEs for the 3D vector velocity \mathbf{u} , vector magnetic field \mathbf{B} , pressure perturbation P' (perturbation away from the “reference-state” pressure \tilde{P}), and entropy perturbation S' (perturbation away from \tilde{S}). Note that S' can also be interpreted as a temperature perturbation in Boussinesq mode. For more details, see **Rayleigh**’s [Documentation](#).

We use spherical coordinates: r (spherical radius), θ (colatitude), ϕ (azimuth angle), as well as cylindrical coordinates: $\lambda = r \sin \theta$ (cylindrical radius, or moment arm) and $z = r \cos \theta$ (axial coordinate). In general, $\hat{\mathbf{e}}_q$ denotes a position-dependent unit vector in the direction of increasing q . With this notation, the full PDE-set solved by **Rayleigh** is:

$$\nabla \cdot (f_1 \mathbf{u}) = 0, \tag{2.1}$$

$$\nabla \cdot \mathbf{B} = 0, \tag{2.2}$$

$$f_1 \left(\frac{D\mathbf{u}}{Dt} + c_1 \hat{\mathbf{e}}_z \times \mathbf{u} \right) = c_2 f_2 S \hat{\mathbf{e}}_r - c_3 f_1 \nabla \left(\frac{P}{f_1} \right),$$

$$+ c_4 (\nabla \times \mathbf{B}) \times \mathbf{B} + c_5 \nabla \cdot \mathbf{D}, \quad (2.3a)$$

$$\text{where} \quad D_{ij} := 2f_1 f_3 \left[e_{ij} - \frac{1}{3} (\nabla \cdot \mathbf{u}) \delta_{ij} \right] \quad (2.3b)$$

$$\text{and} \quad e_{ij} := \frac{1}{2} \left(\frac{\partial u_i}{\partial x_j} + \frac{\partial u_j}{\partial x_i} \right), \quad (2.3c)$$

$$f_1 f_4 \left(\frac{DS}{Dt} + c_{11} f_{14} u_r \right) = c_6 \nabla \cdot (f_1 f_4 f_5 \nabla S)$$

$$+ c_{10} f_6(r) + c_8 c_5 D_{ij} e_{ij} + c_9 c_7 f_7 |\nabla \times \mathbf{B}|^2, \quad (2.4)$$

$$\text{and} \quad \frac{\partial \mathbf{B}}{\partial t} = \nabla \times [\mathbf{u} \times \mathbf{B} - c_7 f_7 \nabla \times \mathbf{B}], \quad (2.5)$$

where $D/Dt := \partial/\partial t + \mathbf{u} \cdot \nabla$ denotes the material derivative.

The shell geometry (i.e., aspect ratio) and reference state fully define the problem. (I have always found that initial conditions seem to ultimately be irrelevant to the final state of the system, although this of course doesn't rule out hysteresis effects from occurring in future cases). The reference state is defined by the spherically-symmetric, time-independent functions $f_i = f_i(r)$ and constants c_j . By adjusting the f_i and c_j , the user can choose between a Boussinesq or anelastic approximation (the specific form of each approximation is described more fully in the [Documentation](#)), choose any arbitrary nondimensionalization of these two equation sets, and/or set the nondimensional parameters (Rayleigh number, Prandtl number, etc.) of the problem. `Rayleigh` has built-in modes to set the f_i and c_j for the single-layer (i.e., either convectively stable or unstable, but not both) Boussinesq approximation (nondimensional only) and anelastic approximation (either dimensional or two choices of nondimensional). These modes are chosen via the choices `reference_type = 1, 2, 3` or `5`.

More complex systems, which `Rayleigh` also supports, require the user to manually change the f_i and c_j . This can be done by editing an input binary file that `Rayleigh` reads upon initialization and setting `reference_type = 4`. This document essentially describes how to produce such a “custom” input file (here called `customfile`) for a coupled stable–unstable anelastic system, using a variety of chosen time-scales in the nondimensionalization. I also unpack what the “energy flux” means when stable layers are present.

Note that the c_j can also be over-written at run-time by an ASCII text-file (i.e., the `main_input` file), allowing easy changes of the nondimensional numbers for a simulation suite that uses a common reference state—all without modifying the `customfile`.

3 Dimensional Anelastic Equations

We begin by writing down the full dimensional anelastic fluid equations, as they are usually implemented in `Rayleigh` (more precisely, this corresponds to `reference_type = 2`). This

form of the anelastic approximation in a spherical shell is derived in, or more accurately, attributed to (since Rayleigh “updates” the reference state slightly differently than the cluge-y ASH implementation), two common sources: Gilman & Glatzmaier (1981) and Clune et al. (1999). Rayleigh’s dimensional anelastic equation-set is:

$$\nabla \cdot (\tilde{\rho} \mathbf{u}) = 0, \quad (3.1)$$

$$\nabla \cdot \mathbf{B} = 0, \quad (3.2)$$

$$\begin{aligned} \tilde{\rho} \left(\frac{D\mathbf{u}}{Dt} + 2\Omega_0 \hat{\mathbf{e}}_z \times \mathbf{u} \right) &= \left(\frac{\tilde{\rho} \tilde{g}}{C_p} \right) S' \hat{\mathbf{e}}_r - \tilde{\rho} \nabla \left(\frac{P'}{\tilde{\rho}} \right), \\ &+ \frac{1}{\mu} (\nabla \times \mathbf{B}) \times \mathbf{B} + \nabla \cdot \mathbf{D}, \end{aligned} \quad (3.3a)$$

$$\text{where} \quad D_{ij} := 2\tilde{\rho} \tilde{\nu} \left[e_{ij} - \frac{1}{3} (\nabla \cdot \mathbf{u}) \delta_{ij} \right] \quad (3.3b)$$

$$\text{and} \quad e_{ij} := \frac{1}{2} \left(\frac{\partial u_i}{\partial x_j} + \frac{\partial u_j}{\partial x_i} \right), \quad (3.3c)$$

$$\tilde{\rho} \tilde{T} \left(\frac{DS'}{Dt} + \frac{d\tilde{S}}{dr} u_r \right) = \nabla \cdot (\tilde{\rho} \tilde{T} \tilde{\kappa} \nabla S') + \tilde{Q}_{\text{Ra}} + D_{ij} e_{ij} + \frac{\tilde{\eta}}{\mu} |\nabla \times \mathbf{B}|^2, \quad (3.4)$$

$$\text{and} \quad \frac{\partial \mathbf{B}}{\partial t} = \nabla \times (\mathbf{u} \times \mathbf{B} - \tilde{\eta} \nabla \times \mathbf{B}). \quad (3.5)$$

Here, the thermal variables ρ , T , P , and S refer to the density, temperature, pressure, and entropy (respectively). The tildes denote the reference state and the primes denote the (assumed small) perturbations from the reference state. More specifically, the essence of the anelastic approximation involves assuming all relative thermal perturbations (ρ'/ρ_a , T'/T_a , P'/P_a , and S'/C_p) are $O(\epsilon)$, where $\epsilon \ll 1$ (e.g., Ogura & Phillips 1962; Gough 1969 and Matilsky & Brummell 2024, coming soon!). Here, the “a” subscripts denote “typical reference-state values” and ϵ is also the same (small) size as the typical squared Mach number of the flow.

Other background quantities that appear are the full gravitational acceleration \tilde{g} , the momentum, thermal, and magnetic diffusivities ($\tilde{\nu}$, $\tilde{\kappa}$, and $\tilde{\eta}$, respectively; these all have c.g.s. units of $\text{cm}^2 \text{s}^{-1}$), Rayleigh’s internal heating \tilde{Q}_{Ra} (much more on this below), the frame rotation rate Ω_0 , the specific heat at constant pressure C_p , and the vacuum permeability μ ($= 4\pi$ in c.g.s. units). The equations are written in a frame rotating with angular velocity $\Omega_0 \hat{\mathbf{e}}_z$ and the centrifugal force is neglected. Note that C_p is always assumed to be constant, while in a real stellar (or gas-giant) structure model, it should technically vary in radius as the “perfectness” of the gas varies.

4 My reference state: a flimsy approach to stellar structure

In the anelastic equations, the one-dimensional reference state should satisfy equilibrium conditions in the momentum and energy equations. These one-dimensional equations are then subtracted from the fully compressible equations and linearized in the primed thermodynamic quantities to produce the anelastic approximation. Ideally, the reference state would be chosen based on a state-of-the-art structure model, incorporating all known physics as best we can. In the Sun, for example, this structure model is fairly well agreed-upon (e.g., [Christensen-Dalsgaard et al. 1996](#)); but in the Jovian case, much less so (e.g., [Guillot 2005](#)). In our (or possibly just Loren’s) philosophy, the form of the reference state itself matters far less than the fact that we cannot approach realistic astrophysical conditions for stars or gas giants. Therefore, we choose the reference state somewhat arbitrarily. Put another way, we choose the f_i to be slightly wrong and are consoled by the fact that the c_j are *incredibly* wrong.

When the reference state is fixed in time as it is in **Rayleigh**, the anelastic equations have no idea whether the one-dimensional equations subtracted out make any sense. It is up to us to be wise and enforce things like hydrostatic balance and thermal equilibrium. **Rayleigh** will happily chug along under any arbitrary reference state with no (obvious) poor consequences. It also helps to recall exactly what structure model we are “really” subtracting off, since there can be a lot of confusion, especially when interpreting “flux through the system.” Clarification of this energy flux is the main purpose of this and the following sections.

To specify our reference state, we demand that the gas be in hydrostatic balance,

$$\frac{d\tilde{P}}{dr} = -\tilde{\rho}\tilde{g}, \quad (4.1)$$

be perfect,

$$\tilde{P} = \tilde{\rho}\mathcal{R}\tilde{T}, \quad (4.2)$$

and be in local thermodynamic equilibrium (LTE),

$$\frac{1}{C_p} \frac{d\tilde{S}}{dr} = \frac{1}{\gamma} \frac{d \ln \tilde{T}}{dr} - \frac{\gamma - 1}{\gamma} \frac{d \ln \tilde{\rho}}{dr} \quad (4.3a)$$

$$= \frac{1}{\gamma} \frac{d \ln \tilde{P}}{dr} - \frac{d \ln \tilde{\rho}}{dr} \quad (4.3b)$$

$$= \frac{d \ln \tilde{T}}{dr} - \frac{\gamma - 1}{\gamma} \frac{d \ln \tilde{P}}{dr} \quad (4.3c)$$

Here, $\mathcal{R} = C_p - C_v$ is the gas constant, C_v is the specific heat at constant volume (or equivalently, constant density), and $\gamma = C_p/C_v$ is the ratio of specific heats. All are assumed to be constant (that is the definition of a perfect, not just ideal, gas).

Really, Equation (4.3) is simply an expression of the first law of thermodynamics. It doesn't further constrain the structure problem, but rather introduces a new unknown, i.e., the entropy \tilde{S} . The approach I employ is simply to choose $d\tilde{S}/dr$ to enforce convective stability or not and also pick $\tilde{g} \propto 1/r^2$. Then the structure problem is easily solved!

5 The real structure problem

In an ideal world, we would not choose an arbitrary $d\tilde{S}/dr$, but instead demand a condition for thermal (in stars, “radiative”) equilibrium. This would generate another equation constraining $\tilde{\rho}$ and \tilde{T} in addition to Equations (4.1)–(4.3), and thus fully define the structure problem. Let's consider radiative equilibrium first, since this applies in many stellar interiors and (maybe) in the weather layer of Jupiter. The radiative diffusion approximation in stars makes the structure problem much easier than for the (relatively cold) gas giants, since we only need consider the opacity of a fully ionized gas (and even then, it's complicated). Under the radiative diffusion approximation (and making a mixing-length assumption), the equation of radiative equilibrium is

$$\tilde{Q} - \nabla \cdot \tilde{\mathbf{F}}_{\text{rad}} = \nabla \cdot \tilde{\mathbf{F}}_{\text{conv}}, \quad (5.1a)$$

$$\text{where} \quad \tilde{\mathbf{F}}_{\text{rad}} := -\tilde{\rho} C_p \tilde{\kappa}_{\text{rad}} \nabla \tilde{T} = -\tilde{\rho} C_p \tilde{\kappa}_{\text{rad}} \frac{d\tilde{T}}{dr} \hat{\mathbf{e}}_r, \quad (5.1b)$$

$$\nabla \cdot \tilde{\mathbf{F}}_{\text{conv}} := \begin{cases} 0 & \text{for } \nabla_{\text{rad}} < \nabla_{\text{ad}} \text{ (stable to convection)} \\ [\tilde{Q} - \nabla \cdot \tilde{\mathbf{F}}_{\text{rad}}]_{(\tilde{\rho}, \tilde{T})=(\tilde{\rho}_{\text{ad}}, \tilde{T}_{\text{ad}})} & \text{for } \nabla_{\text{rad}} > \nabla_{\text{ad}} \text{ (unstable to convection)} \end{cases}, \quad (5.1c)$$

$$\nabla_{\text{rad}} := \frac{d \ln \tilde{T}_{\text{rad}} / dr}{d \ln \tilde{P} / dr} := \frac{\gamma - 1}{\gamma} \frac{\tilde{L}(r)}{4\pi r^2 \tilde{g} \tilde{\rho} \tilde{\kappa}_{\text{rad}}}, \quad (5.1d)$$

$$\text{and} \quad \nabla_{\text{ad}} := \left(\frac{d \ln \tilde{T}}{d \ln \tilde{P}} \right)_{\text{ad}} = \frac{\gamma - 1}{\gamma}. \quad (5.1e)$$

Here, $\tilde{Q} = \tilde{Q}(\tilde{\rho}, \tilde{T})$ is the “real” reference-state internal heating per unit volume (likely due to nuclear burning in the core, or in Jupiter, Helmholtz contraction and solar irradiance), $\tilde{L}(r) \equiv 4\pi \int_0^r \tilde{Q}(x) x^2 dx$ is the total luminosity (power output) of the object interior to r , and $\tilde{\kappa}_{\text{rad}} = \tilde{\kappa}_{\text{rad}}(\tilde{\rho}, \tilde{T})$ is the reference-state radiative diffusivity (units $\text{cm}^2 \text{s}^{-1}$), which in turn depends on the opacity. The notation $(\tilde{\rho}_{\text{ad}}, \tilde{T}_{\text{ad}})$ means “the continuation of $\tilde{\rho}$ and \tilde{T} assuming adiabatic stratification.”

In practice, Equations (4.1)–(5.1) would be solved by integrating inward from an appropriate boundary condition at the photosphere. At each point, the relative values of ∇_{rad} and ∇_{ad} would be compared to determine which of the two forms of Equation (5.1c) to use. Doing so implicitly adheres to the mixing-length assumption, because it is implied that “convection is locally determined.” Wherever $\nabla_{\text{rad}} < \nabla_{\text{ad}}$ there is no convection at all ($\tilde{\mathbf{F}}_{\text{conv}} = 0$) and wherever $\nabla_{\text{rad}} > \nabla_{\text{ad}}$ the convection is so vigorous that the layer is completely adiabatically stratified. No overshoot is allowed!

6 Energy fluxes: what belongs to the reference state?

The reference-state \tilde{Q} and the \tilde{Q}_{Ra} used by [Rayleigh](#) have physically quite different origins. Really, \tilde{Q}_{Ra} is shorthand for $\nabla \cdot \tilde{\mathbf{F}}_{\text{conv}}$. We describe explicitly how this works here.

Subtracting Equation (5.1) from the fully compressible heat equation (see [Matilsky & Brummell 2024](#)) yields

$$\tilde{\rho} \tilde{T} \left(\frac{DS'}{Dt} + \frac{d\tilde{S}}{dr} u_r \right) = Q' - \nabla \cdot \mathbf{F}'_{\text{rad}} + \nabla \cdot \tilde{\mathbf{F}}_{\text{conv}} + D_{ij} e_{ij} + \frac{\tilde{\eta}}{\mu} |\nabla \times \mathbf{B}|^2, \quad (6.1a)$$

$$\text{where} \quad \mathbf{F}'_{\text{rad}} := -\tilde{\rho} C_p \tilde{\kappa}_{\text{rad}} \nabla T' - \left(\frac{\rho'}{\tilde{\rho}} + \frac{\kappa'_{\text{rad}}}{\tilde{\kappa}_{\text{rad}}} \right) \tilde{\rho} C_p \tilde{\kappa}_{\text{rad}} \nabla \tilde{T} \quad (6.1b)$$

in place of Equation (3.4).

In [Rayleigh](#) therefore, we have implicitly set $Q' \equiv 0$ (i.e., we ignore the “real” internal heating processes; so we also set $\tilde{Q} \equiv 0$) and we choose

$$\tilde{Q}_{\text{Ra}} := \nabla \cdot \tilde{\mathbf{F}}_{\text{conv}} = -\nabla \cdot \tilde{\mathbf{F}}_{\text{rad}} \quad (\text{since } \tilde{Q} \equiv 0). \quad (6.2)$$

However, we don’t pick $\tilde{\mathbf{F}}_{\text{rad}}$ from a structure model, rather we set it proportional to \tilde{P} . In the Sun, this isn’t too bad an assumption actually (see [Featherstone & Hindman 2016](#) and the appendix of [Matilsky et al. 2024](#)).

We further implicitly assume

$$\mathbf{F}'_{\text{rad}} = -\tilde{\rho} \tilde{T} \tilde{\kappa} \nabla S'. \quad (6.3)$$

This assumption appears to originate in [Gilman & Glatzmaier \(1981\)](#), which assumes a turbulent “eddy” convective heat flux, due to unresolved motions, of the form $-\rho T \kappa \nabla S$ in the fully compressible heat equation. Making the choice to diffuse entropy, we must therefore regard κ (and thus $\tilde{\kappa}$) as an “eddy” diffusivity. Strictly, we really shouldn’t keep having $\mathbf{F}'_{\text{rad}} \propto \nabla S'$ in radiative zones, where there are no “eddies” (although maybe there are two-dimensional eddies and also secondary vertical shear instabilities that can mix things—see [Cope et al. 2020](#); [Garaud 2020](#)). Hopefully the baroclinic consequences of keeping $\mathbf{F}'_{\text{rad}} \propto \nabla S'$ in radiative zones isn’t so bad!

Why all this rigamarole? It clarifies several things. First, it really only makes sense to add “code internal heating” (like \tilde{Q}_{Ra}) to the *convective* portions of the simulation. Really, $\tilde{Q}_{\text{Ra}} = \nabla \cdot \tilde{\mathbf{F}}_{\text{conv}}$ and it only happens to look like $-\nabla \cdot \tilde{\mathbf{F}}_{\text{rad}}$ in our case because we exclude the “real” internal heating \tilde{Q} . Furthermore, nonzero conductive flux boundary conditions (i.e., $\partial S'/\partial r \neq 0$), which act like sources and sinks of heat in the same way as \tilde{Q}_{Ra} , should not be added to the boundaries of radiative zones.

Second, the Jupiter irradiance, which heats a stable layer, really does *not* constitute “heating from above” in the same way that the solar radiation, which heats a convection zone, constitutes “heating from below.” This is because a weather layer fully in radiative equilibrium would have $\tilde{Q}_{\text{Ra}} = \nabla \cdot \tilde{\mathbf{F}}_{\text{rad}} \equiv 0$, and there is no unbalanced heating. The only place “heating from above” comes in is in the boundary condition for $\tilde{\mathbf{F}}_{\text{rad}}$, which is affected by the irradiance. This alters the reference state, but it doesn’t heat or cool the fluid.

The situation could (philosophically) be different in the sense that absorption of all wavelengths of solar radiation in the upper layers could *act* like a heating from above (i.e, take the form of a volumetric heating \tilde{Q} , that occupied a particular outer layer independent of the temperature gradients there), while re-emission would occur (maybe in the infra-red) by altering the temperature gradient to enforce $\nabla \cdot \tilde{\mathbf{F}}_{\text{rad}} = \tilde{Q}$. But regardless, this balance is subtracted off the heat equation and again leaves no unbalanced “code internal heating” (i.e., $\tilde{Q}_{\text{Ra}} = \nabla \cdot \tilde{\mathbf{F}}_{\text{conv}}$ still = 0).

Third and finally, when we think of anelastic energy fluxes, these fluxes are only responsible for carrying out L_{KH} , i.e., the Kelvin-Helmholtz contraction luminosity that is internal to Jupiter. There is no flux necessary to account for L_{irrad} (the luminosity associated with energy absorbed through irradiance), because this flux has been subtracted out and should be associated with the reference state. We *do* need to set an outward conductive flux at the top of the weather layer (i.e., $\partial S'/\partial r < 0$) in order to maintain thermal equilibrium in the convecting state. This is unfortunate given the convective origin of entropy diffusion discussed above. Its baroclinic consequences (which now may be relevant, given our goal to investigate “shallow driving”) have not been explored to my knowledge. Note that the top conductive flux will make the weather layer less stable than it otherwise would be. But as long as the weather layer is sufficiently stiff (i.e., has high buoyancy frequency), it should be a minimal effect.

This is all to say, I am now fully converted to the idea of setting stability via $d\tilde{S}/dr$, instead of using heating layers in different places. I also view [Heimpel et al. \(2022\)](#), who implements a stable layer above the convective layer, as effectively using heating layers also, since they set the conductive flux to be “into the layer” at both boundaries. This approach has the added disadvantage that the location between convectively stable and unstable cannot be set a priori.

7 Non-Dimensional Scheme

In this section, H denotes the typical length-scale (chosen to be the CZ thickness) and τ the typical time-scale (left general for now).

We recall the relation,

$$\tilde{N}^2 = \frac{\tilde{g}}{C_p} \frac{d\tilde{S}}{dr}, \quad (7.1)$$

where \tilde{N}^2 is the squared buoyancy frequency, which we will use in favor of $d\tilde{S}/dr$ in subsequent equations.

We define the “energy flux not carried by radiation” via

$$\tilde{F}_{\text{nrad}} := \frac{1}{r^2} \int_{r_{\text{bcz}}}^r \tilde{Q}_{\text{Ra}}(x) x^2 dx, \quad (7.2)$$

where r_{bcz} is the radius of the base of the CZ. We also denote the top of the CZ by r_{tcz} , the base of the RZ by r_{brz} , and the top of the RZ by r_{trz} . Note that \tilde{F}_{nrad} is basically similar to $|\tilde{\mathbf{F}}_{\text{conv}}|$, however in the context of stellar structure theory, $\tilde{\mathbf{F}}_{\text{conv}}$ is computed according to a definite logic, whereas we choose our \tilde{Q}_{Ra} arbitrarily.

ΔS denotes the (*estimated*) entropy difference across the CZ, and thus a typical value of S' in the CZ. In Rayleigh-Bénard-type convection, the true entropy difference is imposed directly. We could do this as well, by fixing S' itself at the inner and outer boundaries. However, instead we typically set the entropy difference indirectly by adding an internal heating \tilde{Q}_{Ra} and balancing it with the conductive energy flux at the boundaries (which we set by fixing $\partial S'/\partial r$ at the boundaries in place of S' itself). We estimate the resultant entropy difference via

$$\Delta S := \frac{F_{\text{nrad,a}} H}{\rho_a T_a \kappa_a}, \quad (7.3)$$

The “a” subscripts again refer to “typical values” of the (dimensional) atmospheric reference-state functions. Equation (7.3) essentially assumes that there is a conductive boundary layer (at the top of the CZ, since \tilde{Q}_{Ra} must be concentrated more toward the base of the CZ in order to drive convection) that carries the full internal luminosity $L := 4\pi \int_{r_{\text{bcz}}}^{r_{\text{tcz}}} \tilde{Q}_{\text{Ra}}(x) x^2 dx$ out of the CZ. Note that this is a crude estimate, since the ultimate boundary layer thickness is unknown (and will be substantially less than H). Furthermore, we will choose CZ volume-averages for the “typical values” of $\tilde{\rho}$, \tilde{F}_{nrad} , etc., which will be higher than the boundary-layer appropriate values. We can easily see how bad our assumption is, by calculating the entropy difference across the CZ post-equilibration and comparing to 1.

We now nondimensionalize Equations (3.1)–(3.5), according to the following scheme:

$$\nabla \rightarrow \frac{1}{H} \nabla, \quad (7.4a)$$

$$t \rightarrow \tau t, \quad (7.4b)$$

$$\mathbf{u} \rightarrow \frac{H}{\tau} \mathbf{u}, \quad (7.4c)$$

$$S' \rightarrow (\Delta S) S', \quad (7.4d)$$

$$P' \rightarrow \tilde{\rho} \frac{H^2}{\tau^2} P', \quad (7.4e)$$

$$\mathbf{B} \rightarrow (\mu \tilde{\rho})^{1/2} \frac{H}{\tau} \mathbf{B}, \quad (7.4f)$$

$$\tilde{\rho} \rightarrow \rho_a \tilde{\rho}, \quad (7.4g)$$

$$\tilde{T} \rightarrow T_a \tilde{T}, \quad (7.4h)$$

$$\tilde{g} \rightarrow g_a \tilde{g}, \quad (7.4i)$$

$$\widetilde{N^2} \rightarrow N_a^2 \widetilde{N^2}, \quad (7.4j)$$

$$\tilde{S} \rightarrow C_p \tilde{S}, \quad (7.4k)$$

$$\tilde{\nu} \rightarrow \nu_a \tilde{\nu}, \quad (7.4l)$$

$$\tilde{\kappa} \rightarrow \kappa_a \tilde{\kappa}, \quad (7.4m)$$

$$\tilde{\eta} \rightarrow \eta_a \tilde{\eta}, \quad (7.4n)$$

$$\text{and } \tilde{Q}_{\text{Ra}} \rightarrow \frac{F_{\text{nrad,a}}}{H} \tilde{Q}_{\text{Ra}} \quad (7.4o)$$

On the right-hand-sides of Equation (7.4) and in the following equations, all fluid variables, coordinates, and background-state quantities are understood to be nondimensional. The

“typical values” denoted by the “a” subscripts are chosen to be volume-averages over the CZ, except for N_a^2 , which is chosen to be the volume-average over the stably stratified RZ (really “weather layer” in the case of Jupiter).

We define the relative thicknesses of the RZ and CZ via the “first aspect ratio”

$$\alpha := \frac{r_{\text{trz}} - r_{\text{brz}}}{r_{\text{tcz}} - r_{\text{bcz}}} \quad (7.5)$$

and the aspect ratio of the CZ (the “second aspect ratio”) via

$$\beta = \frac{r_{\text{bcz}}}{r_{\text{tcz}}}. \quad (7.6)$$

Since H is the thickness of the CZ, we have

$$r_{\text{bcz}} = \frac{\beta}{1 - \beta}, \quad (7.7a)$$

$$r_{\text{tcz}} = \frac{1}{1 - \beta}, \quad (7.7b)$$

and

$$r_{\text{brz}} = r_{\text{trz}} - \alpha, \quad (7.8a)$$

$$\text{with } r_{\text{trz}} = r_{\text{bcz}} \quad \text{for the solar tachocline} \quad (7.8b)$$

and

$$r_{\text{brz}} = r_{\text{tcz}}, \quad (7.9a)$$

$$\text{and } r_{\text{trz}} = r_{\text{brz}} + \alpha \quad \text{for Jupiter.} \quad (7.9b)$$

The only difference between Jupiter and the tachocline is which of the CZ or RZ is on top.

We leave the time-scale general for now, but ultimately assume it is a rotational time-scale when running the simulations [i.e., $\tau = (2\Omega_0)^{-1}$]. Leaving τ general allows easy translation between different nondimensionalizations, if they are needed later.

8 Non-Dimensional Equations (General τ)

There are several time-scales relevant to the problem. Putting the nondimensional parameters in terms of these more clearly elucidates their meaning. Practically, it also makes translating between different nondimensionalizations easier. The relevant (dimensional) timescales

are:

$$\tau_\nu := \frac{H^2}{\nu_a} \quad \text{viscous diffusion time (across the CZ),} \quad (8.1a)$$

$$\tau_\kappa := \frac{H^2}{\kappa_a} \quad \text{thermal diffusion time (across the CZ)} \quad (8.1b)$$

$$\tau_\eta := \frac{H^2}{\eta_a} \quad \text{magnetic diffusion time (across the CZ)} \quad (8.1c)$$

$$\tau_{\text{ff}} := \sqrt{\frac{C_p H}{\Delta S g_a}} \quad \text{“effective” free-fall time,} \quad (8.1d)$$

$$\tau_\Omega := \frac{1}{2\Omega_0} \quad \text{rotational (Coriolis) time-scale,} \quad (8.1e)$$

$$\tau_N := \frac{1}{N_a} \quad \text{buoyancy (gravity-wave) time-scale,} \quad (8.1f)$$

$$\tau_{\text{ES}} := \frac{N_a^2}{4\Omega_0^2} \frac{H^2}{\kappa_a} = \left(\frac{\tau_\Omega^2}{\tau_N^2} \right) \tau_\kappa \quad \text{Eddington-sweet time} \quad (8.1g)$$

The relevant nondimensional parameters are:

$$\text{Ra} := \frac{\tau_\nu \tau_\kappa}{\tau_{\text{ff}}^2} = \frac{g_a H^3}{\nu_a \kappa_a} \frac{\Delta S}{C_p} \quad (\text{Rayleigh number}), \quad (8.2a)$$

$$\text{Pr} := \frac{\tau_\kappa}{\tau_\nu} = \frac{\nu_a}{\kappa_a} \quad (\text{Prandtl number}), \quad (8.2b)$$

$$\text{Pr}_m := \frac{\tau_\eta}{\tau_\nu} = \frac{\nu_a}{\tilde{\eta}} \quad (\text{magnetic Prandtl number}), \quad (8.2c)$$

$$\text{Ek} := \frac{\tau_\Omega}{\tau_\nu} = \frac{\nu_a}{2\Omega_0 H^2} \quad (\text{Ekman number}), \quad (8.2d)$$

$$\text{Bu} := \frac{\tau_\nu \tau_\kappa}{\tau_N^2} = \frac{N_a^2 H^4}{\nu_a \kappa_a} \quad (\text{buoyancy number}), \quad (8.2e)$$

$$\text{and} \quad \text{Di} = \frac{g_a H}{C_p \tilde{T}} \quad (\text{dissipation number}), \quad (8.2f)$$

Note that in our convention, the dissipation number is not an independent control parameter, but a function of the non-dimensional parameters characterizing the reference state (this will be seen in Section ??).

Assuming rotation is present, we also define some derivative nondimensional parameters

that can be formed from the parameters in Equation (8.2):

$$\text{Ra}^* := \frac{\tau_\Omega^2}{\tau_{\text{ff}}^2} = \frac{\text{RaEk}^2}{\text{Pr}} = \frac{g_a}{4\Omega_0^2 H} \frac{\Delta S}{c_p} \quad (\text{modified Rayleigh number}), \quad (8.3a)$$

$$\text{Ro}_c := \frac{\tau_\Omega}{\tau_{\text{ff}}} = \text{Ek} \sqrt{\frac{\text{Ra}}{\text{Pr}}} \quad (\text{convective Rossby number}), \quad (8.3b)$$

$$\text{Bu}^* := \frac{\tau_\Omega^2}{\tau_N^2} = \frac{\tau_{\text{ES}}}{\tau_\kappa} = \frac{\text{BuEk}^2}{\text{Pr}} = \frac{N_a^2}{4\Omega_0^2} \quad (\text{modified buoyancy number}), \quad (8.3c)$$

$$\sigma := \sqrt{\frac{\tau_{\text{ES}}}{\tau_\nu}} = \frac{N_a}{2\Omega_0} \sqrt{\text{Pr}} \quad (\text{"}\sigma\text{" parameter}) \quad (8.3d)$$

After applying the nondimensional scheme described by Equation (7.4), Equations (3.1)–(3.5) become

$$\nabla \cdot (\tilde{\rho} \mathbf{u}) = 0, \quad (8.4)$$

$$\nabla \cdot \mathbf{B} = 0, \quad (8.5)$$

$$\begin{aligned} \tilde{\rho} \left(\frac{D\mathbf{u}}{Dt} + \frac{\tau}{\tau_\Omega} \hat{\mathbf{e}}_z \times \mathbf{u} \right) &= \frac{\text{Ra}}{\text{Pr}} \left(\frac{\tau}{\tau_\nu} \right)^2 \tilde{\rho} \tilde{g} S' \hat{\mathbf{e}}_r - \tilde{\rho} \nabla \left[\frac{P}{\tilde{\rho}} \right], \\ &+ (\nabla \times \mathbf{B}) \times \mathbf{B} + \left(\frac{\tau}{\tau_\nu} \right) \nabla \cdot \mathbf{D}, \end{aligned} \quad (8.6a)$$

$$\text{where} \quad D_{ij} := 2\tilde{\rho}\tilde{\nu} \left[e_{ij} - \frac{1}{3}(\nabla \cdot \mathbf{u})\delta_{ij} \right] \quad (8.6b)$$

$$\text{and} \quad e_{ij} := \frac{1}{2} \left(\frac{\partial u_i}{\partial x_j} + \frac{\partial u_j}{\partial x_i} \right), \quad (8.6c)$$

$$\begin{aligned} \tilde{\rho}\tilde{T} \left(\frac{DS'}{Dt} + \frac{\text{Bu}}{\text{Ra}} \frac{\widetilde{N^2}}{\tilde{g}} u_r \right) &= \frac{1}{\text{Pr}} \left(\frac{\tau}{\tau_\nu} \right) \nabla \cdot (\tilde{\rho}\tilde{T}\tilde{\kappa}\nabla S') \\ &+ \frac{1}{\text{Pr}} \left(\frac{\tau}{\tau_\nu} \right) \tilde{Q} + \frac{\text{PrDi}}{\text{Ra}} \left(\frac{\tau_\nu}{\tau} \right) \left(D_{ij}e_{ij} + \frac{1}{\text{Pr}_m} \tilde{\eta} |\nabla \times \mathbf{B}|^2 \right), \end{aligned} \quad (8.7)$$

$$\text{and} \quad \frac{\partial \mathbf{B}}{\partial t} = \nabla \times \left[\mathbf{u} \times \mathbf{B} - \frac{1}{\text{Pr}_m} \left(\frac{\tau}{\tau_\nu} \right) \tilde{\eta} \nabla \times \mathbf{B} \right]. \quad (8.8)$$

Note that we have chosen the entropy-scale (and thus the Rayleigh number) based on the internal heating:

$$\Delta S := \frac{\tilde{Q}\tau}{\tilde{\rho}\tilde{T}} = \frac{\tilde{Q}H^2}{\tilde{\rho}\tilde{T}\kappa_a} \quad (8.9a)$$

$$\text{and} \quad \text{Ra} := \frac{g_a \tilde{Q} H^3 \tau}{\tilde{\rho}\tilde{T} c_p \nu_a \kappa_a} = \frac{g_a \tilde{Q} H^5}{\tilde{\rho}\tilde{T} c_p \nu_a \kappa_a^2}, \quad (8.9b)$$

$$(8.9c)$$

where the first equality in each equation is general (it holds for any choice of time-scale τ) and the second equality is specific to the non-rotating case. Essentially, we have assumed that the heating (or cooling) operates on the time-scale τ before the fluid parcel buoyantly moves to another part of the shell, carrying with it an entropy perturbation ΔS (in the non-rotating case, this should happen on the thermal dissipation time-scale $\tau = H^2/\kappa_a$).

The user is thus free to choose the shapes of Q and C , but not their amplitude, since they must have unity volume-averages over the CZ.

The buoyancy number Bu is the ratio of the typical squared buoyancy frequency to the thermal and viscous diffusion times. It is essentially a “second (stable) Rayleigh number”, and measures the stiffness of the stable layer (recall N_a^2 refers to the typical value of \tilde{N}^2 in the WL). The buoyancy number is independent of the Rayleigh number, which estimates the ultimate instability of the CZ.

9 Non-Dimensional Equations, Rotating [$\tau = (2\Omega_0)^{-1}$]

In the previous section, t (and things with time in the dimensions) was implied to mean $(\kappa_a/H^2)t_{\text{dim}}$, where t_{dim} was the dimensional time. We now want to use a new non-dimensional time, $t_{\text{new}} = \Omega_0 t_{\text{dim}} = (\text{Pr}/\text{Ek})t$. We can thus find the new equations easily from Equations (??)–(8.8). Every place we see a time dimension, we recall $t = (\text{Ek}/\text{Pr})t_{\text{new}}$, so we multiply the place where the time-dimension appears by (Ek/Pr) and drop the “new” subscript [e.g., $t \rightarrow (\text{Ek}/\text{Pr}) t$, $\mathbf{u} \rightarrow (\text{Pr}/\text{Ek})\mathbf{u}$, etc.]

Note that we should now choose a different typical entropy-scale and corresponding Rayleigh number:

$$\Delta S^* := \frac{\tilde{Q}\tau}{\tilde{\rho}\tilde{T}} = \frac{\tilde{Q}}{2\Omega_0\tilde{\rho}\tilde{T}} \quad (9.1a)$$

$$\text{and} \quad \text{Ra} := \frac{g_a\tilde{Q}H^3\tau}{\tilde{\rho}\tilde{T}c_p\nu_a\kappa_a} = \frac{g_a\tilde{Q}H^3}{2\Omega_0\tilde{\rho}\tilde{T}c_p\nu_a\kappa_a}. \quad (9.1b)$$

The reasoning here is that under the influence of rapid rotation, the life-times of upflows or downflows are no longer set by the thermal dissipation time, but by the rotation period. Thus, the heating or cooling of a fluid parcel occurs on a shorter time-scale, leading to a smaller entropy difference across the shell than in the non-rotating case with the same amount of heating. Of course, neither Equations (8.9a) or (9.1a) are particularly convincing estimates and there is a large degree of uncertainty in the actual magnitude of typical entropy perturbations. We can only see how good these estimates are (after the fact) by checking if the achieved (non-dimensional) entropy difference across the shell winds up being close to unity.

We thus find (after rearranging terms),

$$\nabla \cdot [\tilde{\rho}\mathbf{u}] = 0, \quad (9.2)$$

$$\nabla \cdot \mathbf{B} = 0, \quad (9.3)$$

$$\begin{aligned} \tilde{\rho} \left[\frac{D\mathbf{u}}{Dt} + \hat{\mathbf{e}}_z \times \mathbf{u} \right] &= \text{Ra}^* \tilde{\rho} \tilde{g} S \hat{\mathbf{e}}_r - \tilde{\rho} \nabla \left[\frac{P}{\tilde{\rho}} \right], \\ &+ (\nabla \times \mathbf{B}) \times \mathbf{B} + \text{Ek} \nabla \cdot \mathbf{D}, \end{aligned} \quad (9.4a)$$

$$\text{where} \quad D_{ij} := 2\tilde{\rho}\tilde{\nu} \left[e_{ij} - \frac{1}{3}(\nabla \cdot \mathbf{u})\delta_{ij} \right] \quad (9.4b)$$

$$\text{and} \quad e_{ij} := \frac{1}{2} \left(\frac{\partial u_i}{\partial x_j} + \frac{\partial u_j}{\partial x_i} \right), \quad (9.4c)$$

$$\begin{aligned} \tilde{\rho}\tilde{T} \left[\frac{DS}{Dt} + \frac{\text{Bu}^*}{\text{Ra}^*} \frac{\widetilde{N^2}}{\tilde{g}} u_r \right] &= \frac{\text{Ek}}{\text{Pr}} \nabla \cdot [\tilde{\rho}\tilde{T}\tilde{\kappa}\nabla S] \\ &+ Q - C + \frac{\text{DiEk}}{\text{Ra}^*} D_{ij}e_{ij} + \frac{\text{DiEk}}{\text{Pr}_m \text{Ra}^*} \tilde{\eta} |\nabla \times \mathbf{B}|^2, \end{aligned} \quad (9.5)$$

$$\text{and} \quad \frac{\partial \mathbf{B}}{\partial t} = \nabla \times \left[\mathbf{u} \times \mathbf{B} - \frac{\text{Ek}}{\text{Pr}_m} \eta \nabla \times \mathbf{B} \right], \quad (9.6)$$

where we have introduced two modified non-dimensional numbers:

$$\text{Ra}^* := \frac{\text{Ek}^2}{\text{Pr}} \text{Ra} = \frac{g_a}{H\Omega_0^2} \frac{\Delta S}{c_p} = \frac{g_a \tilde{Q}}{(2\Omega_0)^3 \tilde{\rho}\tilde{T}c_p H}, \quad (9.7a)$$

$$\text{and} \quad \text{Bu}^* := \frac{\text{Ek}^2}{\text{Pr}} \text{Bu} = \frac{N_a^2}{4\Omega_0^2} \sim \frac{g_a}{H\Omega_0^2} = \frac{1}{\text{geometric oblateness}}. \quad (9.7b)$$

Note that although the “ $d\tilde{S}/dr$ -terms” in the non-dimensionalizations have seemingly different definitions, they are similar, since

$$\frac{\text{Bu}}{\text{Ra}} \sim \frac{c_p}{\Delta S} \quad \text{and} \quad \frac{\text{Bu}^*}{\text{Ra}^*} \sim \frac{c_p}{\Delta S^*}. \quad (9.8)$$

The only difference is in the different estimates ΔS and ΔS^* .

In terms of Rayleigh’s f ’s and c ’s, we compare Equations (9.2)–(9.6) to (2.1)–(2.5) and find:

$$\begin{array}{ll} f_1 \rightarrow \tilde{\rho} & c_1 \rightarrow 1 \\ f_2 \rightarrow \tilde{\rho}\tilde{g} & c_2 \rightarrow \text{Ra}^* \\ f_3 \rightarrow \tilde{\nu} & c_3 \rightarrow 1 \\ f_4 \rightarrow \tilde{T} & c_4 \rightarrow 4 \\ f_5 \rightarrow \tilde{\kappa} & c_5 \rightarrow \text{Ek} \\ f_6 \rightarrow Q - C & c_6 \rightarrow \text{Ek}/\text{Pr} \\ f_7 \rightarrow \tilde{\eta} & c_7 \rightarrow \text{Ek}/\text{Pr}_m \\ \vdots & c_8 \rightarrow \text{Di}/\text{Ra}^* \\ & c_9 \rightarrow \text{Di}/\text{Ra}^* \\ f_{14} \rightarrow \frac{\widetilde{N^2}}{\tilde{g}} & c_{10} \rightarrow 1 \\ & c_{11} \rightarrow \text{Bu}^*/\text{Ra}^* \end{array}$$

10 Hydrostatic, Ideal-Gas, Jovian Stable–Unstable Layer

To model a background CZ and WL in Jupiter, we consider a spherical shell composed of an ideal, hydrostatic gas extending between inner radius r_{in} and outer radius r_{out} . An assumed transition in convective stability occurs near an intermediate radius r_0 , over width δ . More specifically, we choose quartic matching of the entropy gradient between the two layers:

$$\frac{d\tilde{S}}{dr} = \psi_{\text{WL}}(r; r_0, \delta), \quad (10.1)$$

where

$$\psi_{\text{WL}}(r; r_0, \delta) := \begin{cases} 0 & r \leq r_0 \\ 1 - \left[1 - \left(\frac{r-r_0}{\delta}\right)^2\right]^2 & r_0 < r < r_0 + \delta \\ 1 & r \geq r_0 + \delta. \end{cases} \quad (10.2)$$

We also define

$$\psi_{\text{CZ}}(r; r_0, \delta) := 1 - \psi_{\text{WL}}(r; r_0 - \delta, \delta) = \begin{cases} 1 & r \leq r_0 \\ 1 - \left[1 - \left(\frac{r-r_0}{\delta}\right)^2\right]^2 & r_0 - \delta < r < r_0 \\ 0 & r \geq r_0. \end{cases} \quad (10.3)$$

ψ_{WL} thus “senses” only the WL, and ψ_{CZ} “senses” only the CZ.

With this formulation, the CZ is strictly unstable (really, marginally stable, but becomes unstable from the heating and cooling). This ensures that none of the stable gradient “leaks” into the CZ, as happens with (e.g.) tanh matching. We assume a centrally-concentrated mass so that $\tilde{g} \propto 1/r^2$.

It can then be shown that five non-dimensional parameters fully characterize the shell geometry, $\tilde{\rho}$, and \tilde{T} :

$$\alpha := \frac{r_{\text{out}} - r_0}{r_0 - r_{\text{in}}} \quad (\text{WL-to-CZ aspect ratio}), \quad (10.4a)$$

$$\beta := \frac{r_{\text{in}}}{r_0} \quad (\text{CZ aspect ratio}), \quad (10.4b)$$

$$\gamma := \frac{c_p}{c_v} \quad (\text{specific-heat ratio}), \quad (10.4c)$$

$$\delta \quad (\text{stability transition width}), \quad (10.4d)$$

$$N_\rho := \ln \left[\frac{\tilde{\rho}(r_0)}{\tilde{\rho}(r_{\text{out}})} \right] \quad (\text{number of density scale-heights across CZ}), \quad (10.4e)$$

where c_v is the specific heat at constant volume.

We choose H to be the thickness of the CZ ($r_0 - r_{\text{in}} = 1$ and $r_{\text{out}} - r_0 = \alpha$). Thus,

$$r_{\text{in}} = \frac{1 - \beta}{\beta}, \quad (10.5a)$$

$$r_0 = \frac{1}{1 - \beta}, \quad (10.5b)$$

$$\text{and} \quad r_{\text{out}} = \frac{1}{1 - \beta} + \alpha \quad (10.5c)$$

If we choose $\alpha = 0.25$ and $\beta = 0.9$, then $(r_{\text{in}}, r_0, r_{\text{out}}) = (9, 10, 10.25)$.

With the requirement that the volume-average of \tilde{g} over the CZ be unity, we require

$$\tilde{g} = \left[\frac{1 - \beta^3}{3(1 - \beta)^3} \right] \frac{1}{r^2}. \quad (10.6)$$

It can then be shown from the ideal-gas and hydrostatic conditions (e.g., [Matilsky et al. 2023](#)) that

$$\tilde{T} = e^{\tilde{S}} \left[\tilde{T}(r_0) - \text{Di} \int_{r_0}^r \tilde{g}(x) e^{-\tilde{S}(x)} dx \right]. \quad (10.7)$$

and

$$\tilde{\rho} = \tilde{\rho}(r_0) \exp \left[- \left(\frac{\gamma}{\gamma - 1} \right) \tilde{S} \right] \tilde{T}^{1/(\gamma-1)}, \quad (10.8)$$

where (nastily)

$$\text{Di} := \frac{g_{\text{a}} H}{c_{\text{p}} \tilde{T}} = \frac{3\beta(1 - \beta)^2(1 - e^{-N_{\rho}/n})}{(3\beta/2)(1 - \beta^2)(1 - e^{-N_{\rho}/n}) - (1 - \beta^3)(\beta - e^{-N_{\rho}/n})}, \quad (10.9)$$

$$\tilde{T}(r_0) = \frac{(1 - \beta^3)(1 - \beta)}{(3\beta/2)(1 - \beta^2)(e^{N_{\rho}/n} - 1) - (1 - \beta^3)(\beta e^{N_{\rho}/n} - 1)}, \quad (10.10)$$

$$\text{and } n := \frac{1}{1 - \gamma} \quad (10.11)$$

is the polytropic index of the CZ (note that since $d\tilde{S}/dr \equiv 0$ in the CZ, the stratification of the CZ winds up being an adiabatic polytrope). We have also assumed (without loss of generality) that $\tilde{S}(r_0) = 0$, so that $\tilde{S} \equiv 0$ in the CZ). The constant $\tilde{\rho}(r_0)$ ends up not having an analytical expression but can be easily found by integration over the CZ of Equation (10.8) (and I should maybe just find Di and $\tilde{T}(r_0)$ from integration, too, since I am error-prone).

To follow [Jones et al. \(2011\)](#) [Heimpel et al. \(2022\)](#) somewhat, we assume $n = 2$. Apparently this is a common choice for Jupiter that better approximates its weird equation of state without breaking the ideal gas law. It results in the somewhat strange result $\gamma = 3/2$.

We also simplify our lives and assume constant diffusivities:

$$\tilde{\nu} = \tilde{\kappa} = \tilde{\eta} \equiv 1. \quad (10.12)$$

For the heating, we use the typical **Rayleigh** profile $Q \propto \tilde{\rho} \tilde{T}$, but ensure that the heating is fully contained in the CZ:

$$Q \propto \tilde{\rho} \tilde{T} \psi_{\text{CZ}}(r; r_0, \delta) \quad (10.13)$$

For the cooling, we choose a cooling width δ_{cool} :

$$C \propto \exp \left[\frac{r - r_0}{\delta_{\text{cool}}} \right] \psi_{\text{CZ}}(r; r_0, \delta) \quad (10.14)$$

We thus fully define the following geometry and Jovian-ish reference state:

$$\alpha = 0.25, \tag{10.15a}$$

$$\beta = 0.9, \tag{10.15b}$$

$$\gamma = 3/2, \tag{10.15c}$$

$$\delta = \delta_{\text{cool}} = 0.1 \quad (\text{may mess with these}), \tag{10.15d}$$

$$\text{and} \quad N_\rho = 3. \tag{10.15e}$$

References

- Christensen-Dalsgaard, J., Däppen, W., Ajukov, S. V., et al. 1996, *Science*, 272, 1286, doi: [10.1126/science.272.5266.1286](https://doi.org/10.1126/science.272.5266.1286)
- Clune, T., Elliott, J., Miesch, M., Toomre, J., & Glatzmaier, G. 1999, *Parallel Comp.*, 25, 361, doi: [10.1016/s0167-8191\(99\)00009-5](https://doi.org/10.1016/s0167-8191(99)00009-5)
- Cope, L., Garaud, P., & Caulfield, C. P. 2020, *Journal of Fluid Mechanics*, 903, A1, doi: [10.1017/jfm.2020.600](https://doi.org/10.1017/jfm.2020.600)
- Featherstone, N. A., & Hindman, B. W. 2016, *The Astrophysical Journal*, 818, 32, doi: [10.3847/0004-637x/818/1/32](https://doi.org/10.3847/0004-637x/818/1/32)
- Garaud, P. 2020, *The Astrophysical Journal*, 901, 146, doi: [10.3847/1538-4357/ab9c99](https://doi.org/10.3847/1538-4357/ab9c99)
- Gilman, P. A., & Glatzmaier, G. A. 1981, *The Astrophysical Journal Supplement Series*, 45, 335, doi: [10.1086/190714](https://doi.org/10.1086/190714)
- Gough, D. O. 1969, *J. Atmos. Soc.*, 26, 448, doi: [10.1175/1520-0469\(1969\)026<0448:taaftc>2.0.co;2](https://doi.org/10.1175/1520-0469(1969)026<0448:taaftc>2.0.co;2)
- Guillot, T. 2005, *Annual Review of Earth and Planetary Sciences*, 33, 493–530, doi: [10.1146/annurev.earth.32.101802.120325](https://doi.org/10.1146/annurev.earth.32.101802.120325)
- Heimpel, M. H., Yadav, R. K., Featherstone, N. A., & Aurnou, J. M. 2022, *Icarus*, 379, 114942, doi: [10.1016/j.icarus.2022.114942](https://doi.org/10.1016/j.icarus.2022.114942)
- Jones, C., Boronski, P., Brun, A., et al. 2011, *Icar.*, 216, 120, doi: [10.1016/j.icarus.2011.08.014](https://doi.org/10.1016/j.icarus.2011.08.014)
- Matilsky, L. I., & Brummell, N. H. 2024, *Journal of Fluid Mechanics*
- Matilsky, L. I., Brummell, N. H., Hindman, B. W., & Toomre, J. 2023, *The Astrophysical Journal*, doi: [10.48550/ARXIV.2311.10202](https://doi.org/10.48550/ARXIV.2311.10202)
- . 2024, *The Astrophysical Journal*, 962, 189, doi: [10.3847/1538-4357/ad18b2](https://doi.org/10.3847/1538-4357/ad18b2)
- Ogura, Y., & Phillips, N. A. 1962, *Journal of the Atmospheric Sciences*, 19, 173, doi: [10.1175/1520-0469\(1962\)019<0173:saodas>2.0.co;2](https://doi.org/10.1175/1520-0469(1962)019<0173:saodas>2.0.co;2)