

Hydrostatic, ideal-gas reference states in spherical coordinates

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1 Basic assumptions

We wish to derive the thermodynamic state for an ideal, hydrostatic gas in spherical coordinates, assuming a known specific entropy stratification $\bar{S}(r)$ and gravitational acceleration $g(r)$. In keeping with the convention used in, e.g., **Rayleigh** and **ASH**, we will denote the thermal profiles of this reference state using overbars. In mathematical terms, we have:

$$\frac{d\bar{P}}{dr} = -\bar{\rho}(r)g(r) \quad (\text{hydrostatic balance}), \quad (1)$$

$$\text{and} \quad \bar{P}(r) = \bar{\rho}(r)\mathcal{R}\bar{T}(r) \quad (\text{ideal gas law}). \quad (2)$$

Here, $\bar{P}(r)$ is the pressure, $\bar{\rho}(r)$ the density, $\bar{T}(r)$ the temperature, and \mathcal{R} the gas constant. We further assume that the number of degrees of freedom (d.o.f.) is constant throughout, so that \mathcal{R} is independent of radius. We note the relationship between the specific heats for an ideal gas:

$$c_p = \Gamma_1 c_v = \frac{\Gamma_1 \mathcal{R}}{\Gamma_1 - 1}, \quad (3)$$

where $\Gamma_1 := c_p/c_v$.

We note the First Law of Thermodynamics for the radial thermodynamic gradients (easiest starting point is $\bar{T}d\bar{S} = c_v d\bar{T} - (\bar{P}/\bar{\rho}^2)d\bar{\rho}$):

$$\frac{1}{c_v} \frac{d\bar{S}}{dr} = \frac{d \ln \bar{T}}{dr} - (\Gamma_1 - 1) \frac{d \ln \bar{\rho}}{dr}. \quad (4)$$

Differentiating the ideal gas law (2) gives

$$\frac{d \ln \bar{T}}{dr} = \frac{d \ln \rho}{dr} + \frac{d \ln T}{dr}. \quad (5)$$

Combining Equations (4) and (5) yields the following useful forms of the First Law:

$$\frac{1}{c_p} \frac{d\bar{S}}{dr} = \frac{1}{\Gamma_1} \frac{d \ln \bar{T}}{dr} - \frac{\Gamma_1 - 1}{\Gamma_1} \frac{d \ln \bar{\rho}}{dr} \quad (6a)$$

$$= \frac{1}{\Gamma_1} \frac{d \ln \bar{P}}{dr} - \frac{d \ln \bar{\rho}}{dr} \quad (6b)$$

$$= \frac{d \ln \bar{T}}{dr} - \frac{\Gamma_1 - 1}{\Gamma_1} \frac{d \ln \bar{P}}{dr} \quad (6c)$$

Using Equations (3) and (6), we can derive an ordinary differential equation for the temperature:

$$\boxed{\frac{d\bar{T}}{dr} - \left[\frac{1}{c_p} \frac{d\bar{S}}{dr} \right] \bar{T}(r) = -\frac{g(r)}{c_p}.} \quad (7)$$

We can solve Equation (7) via integrating factors, yielding

$$\bar{T}(r) = -\frac{e^{\bar{S}(r)/c_p}}{c_p} \int_{r_0}^r g(x) e^{-\bar{S}(x)/c_p} dx + \bar{T}_0 e^{[\bar{S}(r) - \bar{S}_0]/c_p} \quad (8)$$

Here, we have defined r_0 as an arbitrary radius from which to start the integration. Accordingly, we define the temperature, pressure, density, and entropy at r_0 as \bar{T}_0 , \bar{P}_0 , $\bar{\rho}_0$, and \bar{S}_0 , respectively.

Note that in equation (8), adding a constant σ to the entropy ($\bar{S} \rightarrow \bar{S} + \sigma$) has no effect on the profile $\bar{T}(r)$. Only the relative stratification of entropy is important. We therefore set $\bar{S}_0 := 0$ in what follows. Equation (8) thus becomes

$$\boxed{\bar{T}(r) = e^{\bar{S}(r)/c_p} \left[\bar{T}_0 - \int_{r_0}^r \frac{g(x)}{c_p} e^{-\bar{S}(x)/c_p} dx \right].} \quad (9)$$

With $\bar{T}(r)$ now obtained, we can obtain $\bar{P}(r)$ through integrating the First Law in its form (6c):

$$\boxed{\bar{P}(r) = \bar{P}_0 e^{-\bar{S}(r)/\mathcal{R}} \left[\frac{\bar{T}(r)}{\bar{T}_0} \right]^{\Gamma_1/(\Gamma_1-1)}}. \quad (10)$$

Finally, $\bar{\rho}(r)$ is determined from Equation (2):

$$\boxed{\bar{\rho}(r) = \bar{\rho}_0 e^{-\bar{S}(r)/\mathcal{R}} \left[\frac{\bar{T}(r)}{\bar{T}_0} \right]^{1/(\Gamma_1-1)}}. \quad (11)$$

Thus, the full state of the atmosphere is purely determined by the entropy stratification $d\bar{S}/dr$, the gravitational acceleration $g(r)$, and the two constants \bar{T}_0 and \bar{P}_0 (or \bar{T}_0 and $\bar{\rho}_0$, etc.) It is well-known that with an equation of state, all the thermodynamic variables can be written as functions of the other two. Effectively, we have chosen our two “fundamental variables” to be the entropy and pressure (because of the hydrostatic assumption, we can replace pressure by gravity).

The state specified in equations (9)–(11) is completely general, relying only on the assumptions of a hydrostatic, ideal gas with position-independent specific heats. To move

forward, we must specify $d\bar{S}/dr$ and $g(r)$. For relatively thin convection zones (e.g., the Sun), it is safe to assume that the gravitational acceleration comes purely from a central spherically distributed mass, i.e.,

$$g(r) = \frac{GM_*}{r^2}, \quad (12)$$

where G is the Universal Gravitational Constant and M_* the stellar mass.

2 Adiabatic atmosphere

Considering the mathematical form of (9), the simplest atmosphere to derive is one that is adiabatic (constant-entropy):

$$\bar{S}(r) \equiv 0 \quad \text{and} \quad \frac{d\bar{S}}{dr} \equiv 0. \quad (13)$$

In this case, all the exponentials in (9) cancel to 1 and we find

$$\bar{T}(r) = -\frac{GM_*}{c_p} \int_{r_0}^r \frac{dx}{x^2} + \bar{T}_0$$

or

$$\boxed{\begin{aligned} \bar{T}(r) &= \bar{T}_0 \left[\tilde{a} \left(\frac{r_0}{r} \right) + (1 - \tilde{a}) \right] \quad (\text{adiabatic atmosphere}), \\ \text{where} \quad \tilde{a} &:= \frac{GM_*}{c_p \bar{T}_0 r_0}. \end{aligned}} \quad (14)$$

$$(15)$$

Then, Equations (10) and (11) yield

$$\boxed{\begin{aligned} \bar{P}(r) &= \bar{P}_0 \left[\tilde{a} \left(\frac{r_0}{r} \right) + (1 - \tilde{a}) \right]^{\Gamma_1/(\Gamma_1-1)} \quad (\text{adiabatic atmosphere}) \\ \text{and} \quad \bar{\rho}(r) &= \bar{\rho}_0 \left[\tilde{a} \left(\frac{r_0}{r} \right) + (1 - \tilde{a}) \right]^{1/(\Gamma_1-1)} \quad (\text{adiabatic atmosphere}). \end{aligned}} \quad (16)$$

$$(17)$$

Note: to avoid negative thermodynamic quantities, we must ensure that the argument

$$\tilde{a} \left(\frac{r_0}{r} \right) + (1 - \tilde{a}) > 0. \quad (18)$$

This leads to the conditions on available radii r (for our particular choices of \tilde{a} —i.e., M_* , \mathcal{R} , and \bar{T}_0 —and r_0) for which the derived atmosphere is valid:

$$0 < r < \left(\frac{\tilde{a}}{\tilde{a} - 1} \right) r_0 \quad \tilde{a} > 1 \quad (19a)$$

$$r > 0 \quad \tilde{a} \leq 1. \quad (19b)$$

3 Polytropes

The starting assumption for a “polytrope” is that the temperature gradient is divergenceless—i.e., there is a constant flux due to radiative diffusion throughout the fluid layer. If we assume, for simplicity, a constant radiative diffusion coefficient, then $\nabla^2 \bar{T} = 0$, which, in spherical coordinates, is:

$$\bar{T}(r) = \frac{A}{r} + B. \quad (20)$$

The integration constants A and B in Equation (20) are arbitrary, but one is eliminated from the condition $T(r_0) = \bar{T}_0$, leaving

$$\bar{T}(r) = \bar{T}_0 \left[a \left(\frac{r_0}{r} \right) + (1 - a) \right] \quad (21)$$

where the new constant $a := A/\bar{T}_0 r_0$ is dimensionless. Note the similarity of the argument $[a(r_0/r) + (1 - a)]$ in Equation (21) (which was derived by assuming a divergenceless temperature gradient) to the argument in Equation (14) (which was derived by assuming the gas was adiabatic).

Using Equation (2) to eliminate $\bar{\rho}(r)$ from Equation (1) gives

$$\bar{P}(r) = \bar{P}_0 \exp \left[-\frac{1}{\mathcal{R}} \int_{r_0}^r \frac{g(x)}{\bar{T}(x)} dx \right],$$

from which

$$\bar{P}(r) = \bar{P}_0 \left[a \left(\frac{r_0}{r} \right) + (1 - a) \right]^{n+1}, \quad (22)$$

where we have defined the *polytropic index* n to be

$$n := \frac{b}{a} - 1 \iff a = \frac{b}{n+1}, \quad (23)$$

$$\text{where } b := \frac{GM_*}{\mathcal{R}\bar{T}_0 r_0}. \quad (24)$$

Note that b is closely related to the dissipation number ($\text{Di} := \tilde{g}H/(c_p \tilde{T})$, where H is a typical length-scale and the tildes denote typical background-state values).

We can rewrite the argument in square brackets as:

$$a \left(\frac{r_0}{r} \right) + (1 - a) = \frac{b}{n+1} \left(\frac{r_0}{r} \right) + \frac{n+1-b}{n+1}. \quad (25)$$

Putting it all together (and using the ideal gas law to derive the density), we have

$$\bar{T}(r) = \bar{T}_0 \left[\frac{b}{n+1} \left(\frac{r_0}{r} \right) + \frac{n+1-b}{n+1} \right] \quad (\text{polytropic atmosphere}). \quad (26)$$

$$\bar{P}(r) = \bar{P}_0 \left[\frac{b}{n+1} \left(\frac{r_0}{r} \right) + \frac{n+1-b}{n+1} \right]^{n+1} \quad (\text{polytropic atmosphere}). \quad (27)$$

$$\bar{\rho}(r) = \bar{\rho}_0 \left[\frac{b}{n+1} \left(\frac{r_0}{r} \right) + \frac{n+1-b}{n+1} \right]^n \quad (\text{polytropic atmosphere}). \quad (28)$$

Similarly to the adiabatic atmosphere (see Equation (19)), the choices for a (i.e., b —via M_* , \mathcal{R} , and \bar{T}_0 —and n) and r_0 determine the range of radii for which the polytrope is valid (i.e., the range for which the thermodynamic quantities are positive):

$$0 < r < r_{\max}(n) := \left[\frac{b}{b - (n + 1)} \right] r_0 \quad b = \frac{GM_*}{\mathcal{R}\bar{T}_0 r_0} > n + 1 \quad (29a)$$

$$r > 0 \quad b = \frac{GM_*}{\mathcal{R}\bar{T}_0 r_0} \leq n + 1. \quad (29b)$$

For a particular polytrope (fixed n), the choice of integration origin (r_0 , \bar{T}_0 , and \bar{P}_0) has no effect on the upper bound r_{\max} . However, if $n < b - 1$ (a range on n which *is* affected by the choice of integration origin), the polytropic index n affects the upper bound. Higher n yields a wider range of validity, and lower n a narrower one.

To make things applicable to a real star, we work out where an adiabatic polytrope is valid for the Sun. Recall $G = 6.67 \times 10^{-8} \text{ cm}^3 \text{ g}^{-1} \text{ s}^{-2}$. For the Sun ($M_* = M_\odot = 1.99 \times 10^{33} \text{ g}$), we can integrate from the base of the convection zone ($r_0 = r_{\text{bcz}} = 5.00 \times 10^{10} \text{ cm}$; $\bar{T}_0 = \bar{T}_{\text{bcz}} = 2.11 \times 10^6 \text{ K}$; $c_p = 3.50 \times 10^8 \text{ erg g}^{-1} \text{ K}^{-1}$ or $\mathcal{R} = 1.40 \times 10^8 \text{ erg g}^{-1} \text{ K}^{-1}$) and assume near-adiabatic stratification ($n = 1.50$). Then $b = 8.99 > 2.5$ and condition (29) becomes

$$0 < r < (r_{\max})_\odot := 6.93 \times 10^{10} \text{ cm}.$$

The upper bound is only a little less than the solar radius, $R_\odot = 6.96 \times 10^{10} \text{ cm}$. This makes sense of course, since the solar convection zone *is* (mostly) a polytrope with $n = 1.5$, and the photosphere corresponds to (relatively) very low temperature, density, and pressure.

Note that for the polytrope, in addition to the two integration constants \bar{P}_0 and \bar{T}_0 , there is another constant n (or equivalently, a). Specifying n is the same as specifying $d\bar{S}/dr$. For, plugging Equations (21)–(28) (we only need two of them) into any form of (6) yields

$$\boxed{\begin{aligned} \frac{1}{c_p} \frac{d\bar{S}}{dr} &= \frac{b}{\tilde{n} + 1} \left(\frac{r_0}{r^2} \right) \left[\frac{n - \tilde{n}}{n + 1 - b + b(r_0/r)} \right] \\ \text{and} \quad \frac{\bar{S}(r)}{c_p} &= \left(\frac{n - \tilde{n}}{\tilde{n} + 1} \right) \left\{ \ln \left(\frac{r}{r_0} \right) - \ln \left[\frac{b}{n + 1} + \frac{n + 1 - b}{n + 1} \left(\frac{r}{r_0} \right) \right] \right\} \\ &\quad \text{(polytropic atmosphere),} \end{aligned}} \quad (30a)$$

where

$$\tilde{n} = \frac{1}{\tilde{a}} \frac{GM_*}{\mathcal{R}\bar{T}_0 r_0} - 1 := \frac{1}{\Gamma_1 - 1}. \quad (31)$$

From (30a), it is obvious that $n = \tilde{n}$ (or $a = \tilde{a}$) corresponds to an adiabatic atmosphere (as we have already shown in (14)–(17)), while

$$n > \tilde{n} \iff \frac{d\bar{S}}{dr} > 0 \text{ (atmosphere is stable to convection)} \quad (32)$$

$$\text{and} \quad n < \tilde{n} \iff \frac{d\bar{S}}{dr} < 0 \text{ (atmosphere is unstable to convection)}. \quad (33)$$

Say we want to increase the stability of a solar-like atmosphere by increasing n . What happens? We compute, for a particular radius r :

$$\frac{\partial}{\partial n} \left(\frac{1}{c_p} \frac{d\bar{S}}{dr} \right) = \frac{b}{\tilde{n} + 1} \left(\frac{r_0}{r^2} \right) \left\{ \frac{\tilde{n} + 1 - b + b(r_0/r)}{[n + 1 - b + b(r_0/r)]^2} \right\}. \quad (34)$$

The sign of this partial derivative is independent of n . Therefore, at any given radius r , the magnitude of the entropy gradient monotonically increases or decreases with n . Let's assume we confine ourselves to $r < (r_{\max})_{\odot}$, even though for high n , there will technically be a wider range of validity for the polytrope. Then $\tilde{n} + 1 - b + b(r_0/r) \rightarrow 0$ (from above) as $r \rightarrow (r_{\max})_{\odot}$ (from below). Thus (for sufficiently high n and $r < (r_{\max})_{\odot}$), the partial derivative in Equation (34) is positive, and maximum stability is achieved for $n \rightarrow \infty$.

Recalling that $\lim_{n \rightarrow \infty} (1 + x/n)^n = \exp(x)$, Equations (21)–(30a) show that the limit $n \rightarrow \infty$ simply corresponds to an isothermal atmosphere:

$$\bar{T}(r) \rightarrow \bar{T}_0, \quad (35)$$

$$\bar{P}(r) \rightarrow \bar{P}_0 \exp \left[\frac{GM_*}{\mathcal{R}\bar{T}_0 r_0} \left(\frac{r_0}{r} - 1 \right) \right], \quad (36)$$

$$\bar{\rho}(r) \rightarrow \bar{\rho}_0 \exp \left[\frac{GM_*}{\mathcal{R}\bar{T}_0 r_0} \left(\frac{r_0}{r} - 1 \right) \right], \quad (37)$$

$$\frac{1}{c_p} \frac{d\bar{S}}{dr} \rightarrow \frac{GM_*}{c_p \bar{T}_0 r^2}, \quad (38)$$

$$\text{and } \frac{\bar{S}(r)}{c_p} \rightarrow \frac{GM_*}{c_p \bar{T}_0 r_0} \left(1 - \frac{r_0}{r} \right) \quad (39)$$

(polytropic atmosphere as $n \rightarrow \infty$).

If one tries to increase the stability a solar-like polytrope by increasing n , there is therefore a fundamental upper bound imposed by the polytropic formulation itself. Similarly, as $n \rightarrow \infty$, the total entropy contrast from top to bottom and all logarithmic derivatives of the thermodynamic variables have upper bounds, meaning that all scale heights (which typically set the size of fluid motions) have lower bounds. These bounds are reached for an isothermal atmosphere.

Finally, suppose we want to increase the *instability* of the system by *lowering* n . Note that mathematically, as formulated in Equations (26)–(28), the only value for n *not* allowed is -1 . Physically, $n = -1$ can't be allowed since it would correspond to an isobaric (constant-pressure) atmosphere incapable of hydrostatic support. However, unless we have a star that is denser in the outer layers (is that ever possible? Honest question!), (28) shows we should further restrict ourselves to $n \geq 0$.

As $n \rightarrow 0$, we find the most unstable (isopycnic, or constant-density) atmosphere:

$$\bar{T}(r) \rightarrow \bar{T}_0 \left[b \left(\frac{r_0}{r} \right) + (1 - b) \right], \quad (40)$$

$$\bar{P}(r) \rightarrow \bar{P}_0 \exp \left[\frac{GM_*}{\mathcal{R}\bar{T}_0 r_0} \left(\frac{r_0}{r} - 1 \right) \right], \quad (41)$$

$$\bar{\rho}(r) \rightarrow \bar{\rho}_0 \exp \left[\frac{GM_*}{\mathcal{R}\bar{T}_0 r_0} \left(\frac{r_0}{r} - 1 \right) \right], \quad (42)$$

$$\frac{1}{c_p} \frac{d\bar{S}}{dr} \rightarrow \frac{GM_*}{c_p \bar{T}_0 r^2}, \quad (43)$$

$$\text{and} \quad \frac{\bar{S}(r)}{c_p} \rightarrow \frac{GM_*}{c_p \bar{T}_0 r_0} \left(1 - \frac{r_0}{r} \right) \quad (44)$$

(polytropic atmosphere as $n \rightarrow 0$).

3.1 Polytrope in terms of overall density stratification

For an adiabatic, polytropic convection zone (CZ) sitting atop a stably stratified radiative zone (RZ), it is often convenient to choose the reference radius to be boundary between RZ and CZ, i.e., $r_0 = r_m$ (subscript m for “middle radius”). In the work of [Jones et al. \(2011\)](#) and in the Rayleigh code ([Featherstone 2018](#)), the polytrope is initiated using the middle density ρ_m and the number of density scale heights N_ρ across the CZ, as opposed to the two middle-boundary values p_m and T_m . Here $N_\rho := \ln(\rho_m/\rho_o)$, where ρ_m and ρ_o are the density values at the middle and outer boundaries of the shell, respectively. We can recast (21)–(28) in terms of ρ_m and N_ρ by using (24) and (28) to yield

$$a = \frac{e^{N_\rho/n} - 1}{(1 - \beta)e^{N_\rho/n}} \quad \text{and} \quad T_m = \frac{(1 - \beta)e^{N_\rho/n}}{e^{N_\rho/n} - 1} \frac{GM_*}{(n + 1)\mathcal{R}r_m}, \quad (45)$$

where we have defined the aspect ratio of the CZ,

$$\beta := \frac{r_m}{r_o} < 1. \quad (46)$$

Thus, instead of specifying T_m and p_m , we can specify ρ_m and N_ρ ; T_m is then determined through (45) and p_m through the ideal gas law.

3.2 Polytrope with respect to the center of the CZ

[Jones et al. \(2011\)](#) define the polytrope using the variables at the center of the CZ (i.e., in our notation $r_0 = r_c := (r_m + r_o)/2$). In that case, (21)–(28) become

$$\bar{T}(r) = T_c \left[a_c \left(\frac{r_c}{r} \right) + (1 - a_c) \right], \quad (47)$$

$$\bar{P}(r) = p_c \left[a_c \left(\frac{r_c}{r} \right) + (1 - a_c) \right]^{n+1}, \quad (48)$$

$$\bar{\rho}(r) = \rho_c \left[a_c \left(\frac{r_c}{r} \right) + (1 - a_c) \right]^n, \quad (49)$$

where

$$a_c := \frac{GM_*}{(n+1)\mathcal{R}T_c r_c}. \quad (50)$$

Using a similar computation as the one leading to (45), one can show that

$$a_c = \frac{2\beta(e^{N_\rho/n} - 1)}{(1 - \beta)(\beta e^{N_\rho/n} + 1)}. \quad (51)$$

Defining

$$d := r_o - r_m, \quad (52)$$

$$c_0 := 1 - a_c = \frac{(1 + \beta)(1 - \beta e^{N_\rho/n})}{(1 - \beta)(\beta e^{N_\rho/n} + 1)}, \quad (53)$$

$$c_1 := \left(\frac{r_c}{d}\right) a_c = \frac{(1 + \beta)\beta(e^{N_\rho/n} - 1)}{(1 - \beta)^2(\beta e^{N_\rho/n} + 1)}, \quad (54)$$

$$\text{and} \quad \tilde{\zeta}(r) := a_c \left(\frac{r_c}{r}\right) + (1 - a_c) = c_0 + c_1 \left(\frac{d}{r}\right), \quad (55)$$

(47)–(49) become

$$\bar{T}(r) = T_c \tilde{\zeta}(r), \quad \bar{P}(r) = p_c [\tilde{\zeta}(r)]^{n+1}, \quad \bar{\rho}(r) = \rho_c [\tilde{\zeta}(r)]^n. \quad (56)$$

With significantly more algebra, we can also see that if we define

$$\tilde{\zeta}_o := \tilde{\zeta}(r_o) = \frac{1 + \beta}{\beta e^{N_\rho/n} + 1}, \quad (57)$$

then

$$c_0 = \frac{2\tilde{\zeta}_o - \beta - 1}{1 - \beta} \quad \text{and} \quad c_1 = \frac{(1 + \beta)(1 - \tilde{\zeta}_o)}{(1 - \beta)^2}, \quad (58)$$

yielding exactly the formulation for the polytrope in Jones et al. (2011).

Both the formulation with respect to the center of the shell (55)–(58) and the formulation with respect to the middle boundary using (21)–(28) (also (45) if specifying ρ_m and N_ρ instead of p_m and T_m) are mathematically equivalent. However, given the substantial extra notation involved in defining the constants c_0 and c_1 , it seems preferable to use the latter formulation.

3.3 Stable vs. unstable polytropes

Figure 1 (Figure 2) shows a sample of convectively stable (unstable) polytropes for various values of the polytropic index n in a solar-like convection zone. For very high values of n (Figure 1), the entropy gradient profile asymptotes to its maximum value and the temperature profile becomes flat—in agreement with the discussion surrounding Equation (35). For very low values of n (highly unstable polytropes; Figure 2), the constant a becomes significantly greater than 1, and the argument $\zeta(r)$ becomes negative within the convection zone, giving complex values for $\bar{P}(r)$ and $\bar{\rho}(r)$ for r greater than the radius at which $\zeta = 0$ (the exception to this rule is the special case $n = 1$). Thus, the polytropic formulation is unphysical for domains with large enough extent; the same is true for the stable polytropes, but the radius at which $\zeta = 0$ is quite far outside the convection zone.

3.4 Atmosphere with constant entropy gradient

A simple example of a fluid layer that is stable to convection is one with a constant “stiffness,” or entropy gradient. If we consider an RZ extending from the inner radius r_i of the shell to the middle of the shell r_m , we can write

$$\boxed{\bar{S}(r) = kc_p \left(\frac{r}{r_m} - 1 \right) \quad \text{and} \quad \frac{d\bar{S}}{dr} = \frac{kc_p}{r_m} \quad (\text{constant entropy gradient}),} \quad (59)$$

where k is a dimensionless constant (> 0) representing the stiffness of the stable region. We can then use (9) to find $\bar{T}(r)$:

$$\bar{T}(r) = -\frac{GM_*}{c_p} \left\{ e^{k[(r/r_m)-1]} \right\} \underbrace{\int_{r_m}^r \frac{e^{-k[(r/r_m)-1]} dx}{x^2}}_{:=\mathcal{J}(r)} + T_m e^{k[(r/r_m)-1]}. \quad (60)$$

The integral $\mathcal{J}(r)$ can be recast in terms of the exponential integral function

$$E_n(x) := \int_1^\infty \frac{e^{-xt}}{t^n} dt = x^{n-1} \int_x^\infty \frac{e^{-t}}{t^n} dt, \quad (61)$$

yielding

$$\boxed{\bar{T}(r) = T_m \left\{ \left[e^k \tilde{a} \left(\frac{r_m}{r} \right) E_2 \left(\frac{kr}{r_m} \right) + (1 - e^k E_2(k) \tilde{a}) \right] e^{k[(r/r_m)-1]} \right\},} \quad (62a)$$

$$\boxed{\begin{aligned} \text{with} \quad \tilde{a} &:= \frac{GM_*}{c_p T_m r_m} \quad (\text{again}). \\ &(\text{constant entropy gradient}). \end{aligned}} \quad (62b)$$

Using (10) and (11) then yields

$$\boxed{\begin{aligned} \bar{P}(r) &= p_m \exp \left[-\frac{\gamma}{\gamma-1} k \left(\frac{r}{r_m} - 1 \right) \right] \\ &\times \left\{ \left[e^k \tilde{a} \left(\frac{r_m}{r} \right) E_2 \left(\frac{kr}{r_m} \right) + (1 - e^k E_2(k) \tilde{a}) \right] e^{k[(r/r_m)-1]} \right\}^{\gamma/(\gamma-1)} \end{aligned}} \quad (63)$$

$$\boxed{\begin{aligned} \text{and} \quad \bar{\rho}(r) &= \rho_m \exp \left[-\frac{\gamma}{\gamma-1} k \left(\frac{r}{r_m} - 1 \right) \right] \\ &\times \left\{ \left[e^k \tilde{a} \left(\frac{r_m}{r} \right) E_2 \left(\frac{kr}{r_m} \right) + (1 - e^k E_2(k) \tilde{a}) \right] e^{k[(r/r_m)-1]} \right\}^{1/(\gamma-1)} \\ &(\text{constant entropy gradient}). \end{aligned}} \quad (64)$$

Clearly if $k = 0$, we recover (14)–(17) for an adiabatic atmosphere (note that $E_n(0) \equiv 1$ for all n).

Using (7), one can show that for small stiffnesses (k close to 0), $dT/dr < 0$ throughout the entire RZ, implying a monotonically decreasing temperature profile. However, for $k > k_c$, where

$$k_c := \frac{GM_*}{c_p T_m r_m}, \quad (65)$$

the temperature gradient at $r = r_m$ becomes positive, and there is a temperature inversion somewhere for $r < r_m$. Interestingly, the entropy gradient associated with this transition, $\frac{d\bar{S}}{dr}_c = k_c c_p / r_m = GM_*/T_m r_m^2$ is exactly the same as the entropy gradient at $r = r_m$ for an isothermal atmosphere (or the upper limit on the entropy gradient for a polytropic atmosphere as $n \rightarrow \infty$). For solar values (namely $M = M_\odot$, $r_m = 5.0 \times 10^{10}$ cm, $T_m = 2.1 \times 10^6$ K, and $c_p = 3.5 \times 10^8$ erg g⁻¹ K⁻¹),

$$k_c = 3.59. \quad (66)$$

Figures 3 and 4 show stable fluid layers in a solar-like radiative zone for stable polytropes and constant-entropy fluid layers, respectively. In Figure 3, a range of polytropic indices n are used to more or less give entropy-gradient magnitudes matching those set by the k -values of Figure 4. For $k > 3.59$ (the black and chartreuse curves in Figure 4), the entropy gradients for the constant-gradient layer begin to exceed those of the polytrope, and there is furthermore a noticeable temperature inversion within the radiative zone.

4 Hyperbolic (exponential) matching of entropy gradient between CZ and RZ

We consider a domain in radius of (r_i, R_o) with the middle radius r_m lying in this interval. We simulate a convection zone (CZ) in the region (r_m, R_o) overlying a radiative zone (stable region; RZ) in the region (r_i, r_m) . In order to smoothly match a constant-entropy-gradient RZ ($\frac{d\bar{S}}{dr} = \text{constant} > 0$) to a marginally stable CZ ($\frac{d\bar{S}}{dr} \equiv 0$), we can use hyperbolic trigonometric functions:

$$\bar{S}(r) = \frac{kc_p}{r_m} \left\{ \frac{1}{2} \left[(r - r_m) - \delta \ln \cosh \left(\frac{r - r_m}{\delta} \right) \right] \right\} \quad (67)$$

$$\frac{ds}{dr} = \frac{kc_p}{r_m} \left\{ \frac{1}{2} \left[1 - \tanh \left(\frac{r - r_m}{\delta} \right) \right] \right\}. \quad (68)$$

$$\frac{d^2 s}{dr^2} = -\frac{kc_p}{r_m} \left(\frac{1}{2\delta} \right) \text{sech}^2 \left(\frac{r - r_m}{\delta} \right) \quad (69)$$

We then find the full thermodynamic profiles using the integral relations in (9)–(11).

Figures 5 and 6 show the thermodynamic states defined by the entropy profile (67) for a range of δ -values and k -values. The hyperbolic matching condition has the disadvantage that the layer is slightly non-adiabatic even in the CZ—i.e., the tanh function does not decay quickly enough; indeed, the profile $(1/2)(1 - \tanh(x))$ is still 0.0025 for $x = 3$. From Figure 5, the problem is diminished for very rapid transitions (small δ) but this is also seen to create large spikes in the logarithmic derivatives.

5 Quartic matching of entropy gradient between CZ and RZ

In order to better control precisely where the descending downflow plumes begin to overshoot, we can demand that entropy gradient is *exactly* zero in the CZ, transitioning to a nonzero value (over a width δ) by means of a twice-differentiable piece-wise function for the entropy profile (the entropy gradient $\frac{d\bar{S}}{dr}$ is once-differentiable and has a quartic form):

$$\bar{S}(r) = \begin{cases} \frac{8}{15}k c_p \left(\frac{\delta}{r_m}\right) + k c_p \left(\frac{r}{r_m} - 1\right) & r \leq r_m - \delta \\ k c_p \left(\frac{\delta}{r_m}\right) \left[\frac{2}{3} \left(\frac{r-r_m}{\delta}\right)^3 - \frac{1}{5} \left(\frac{r-r_m}{\delta}\right)^5 \right] & r_m - \delta < r < r_m \\ 0 & r \geq r_m \end{cases} \quad (70)$$

$$\frac{ds}{dr} = \begin{cases} \frac{k c_p}{r_m} & r \leq r_m - \delta \\ \frac{k c_p}{r_m} \left\{ 1 - \left[1 - \left(\frac{r-r_m}{\delta}\right)^2 \right]^2 \right\} & r_m - \delta < r < r_m \\ 0 & r \geq r_m \end{cases} \quad (71)$$

$$\frac{d^2s}{dr^2} = \begin{cases} 0 & r \leq r_m - \delta \\ \frac{4}{\delta} \frac{k c_p}{r_m} \left[1 - \left(\frac{r-r_m}{\delta}\right)^2 \right] \left(\frac{r-r_m}{\delta}\right) & r_m - \delta < r < r_m \\ 0 & r \geq r_m \end{cases} \quad (72)$$

Figures 7 and 8 show the thermodynamic states defined by the entropy profile (70) for a range of δ -values and k -values. We can see that the transition width δ is “strict” in the sense that the entropy gradient is *exactly* equal to its asymptotic values outside the transition region. Thus, the thermodynamic profiles in the CZ are exactly that of an adiabatic polytrope. However, this sharper transition is also seen to create sharper spikes in the logarithmic derivatives.

References

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- Jones, C.A., Boronski, P., Brun, A.S., et al., 2011, *Icarus*, 216, 120

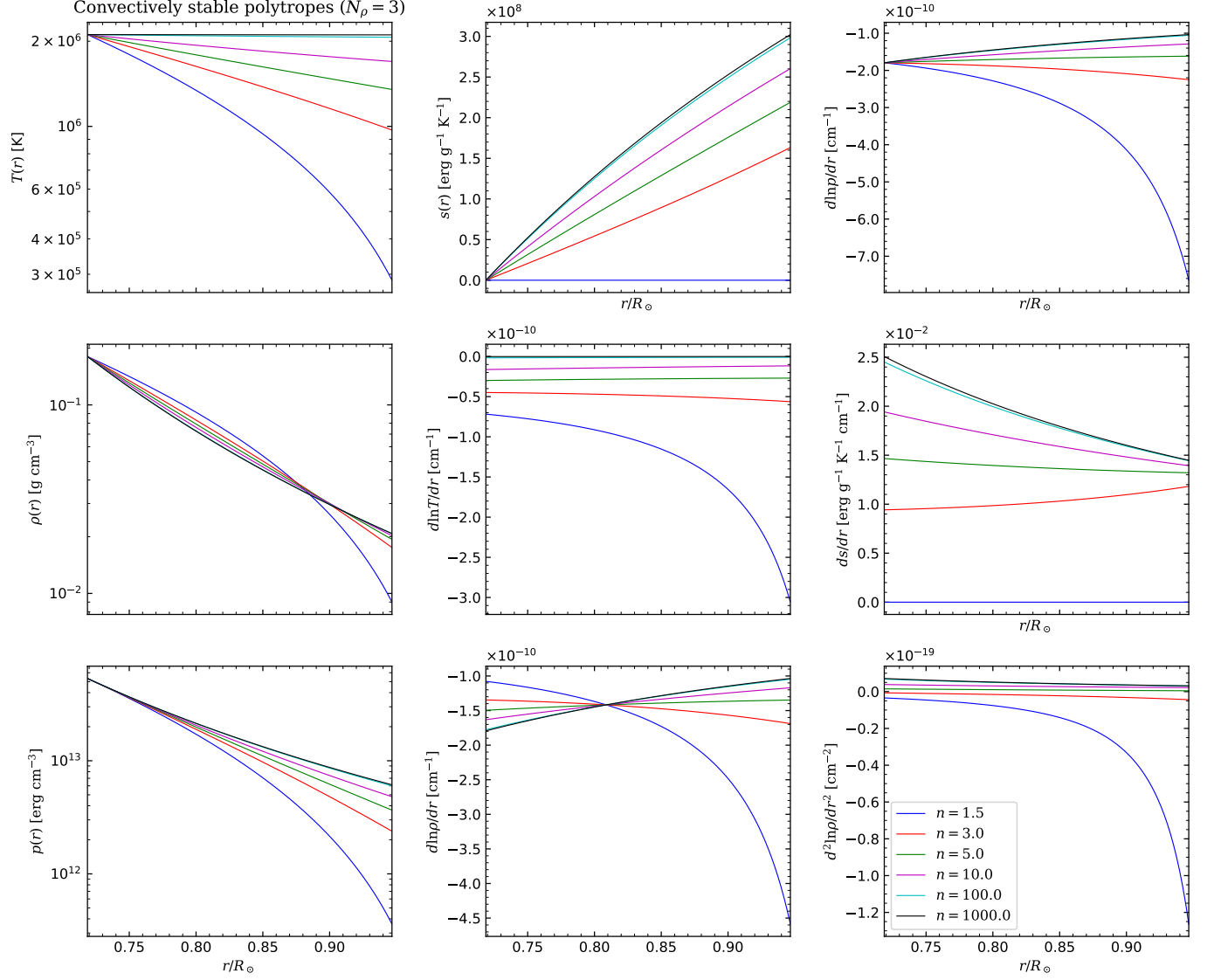


Figure 1: Sample of convectively stable polytropes for a solar-like CZ ($n \geq 1.5$). The dimensional/geometric values for the polytrope (i.e., ρ_i , M , c_p) are the solar values preceding (66), and $N_\rho = 3$.

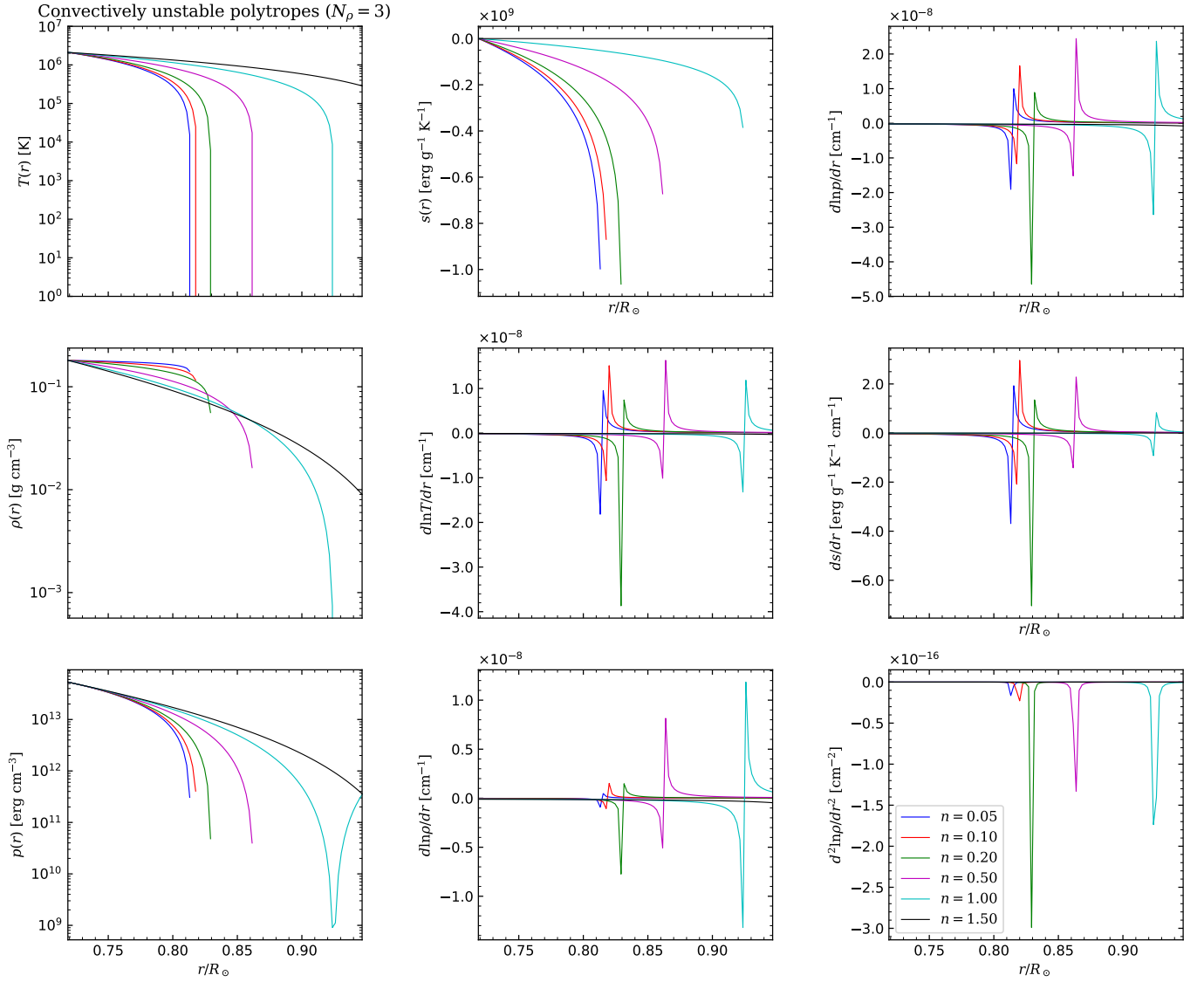


Figure 2: Sample of convectively unstable polytropes for a solar-like CZ ($n \leq 1.5$), with the same dimensional parameters as Figure 1.

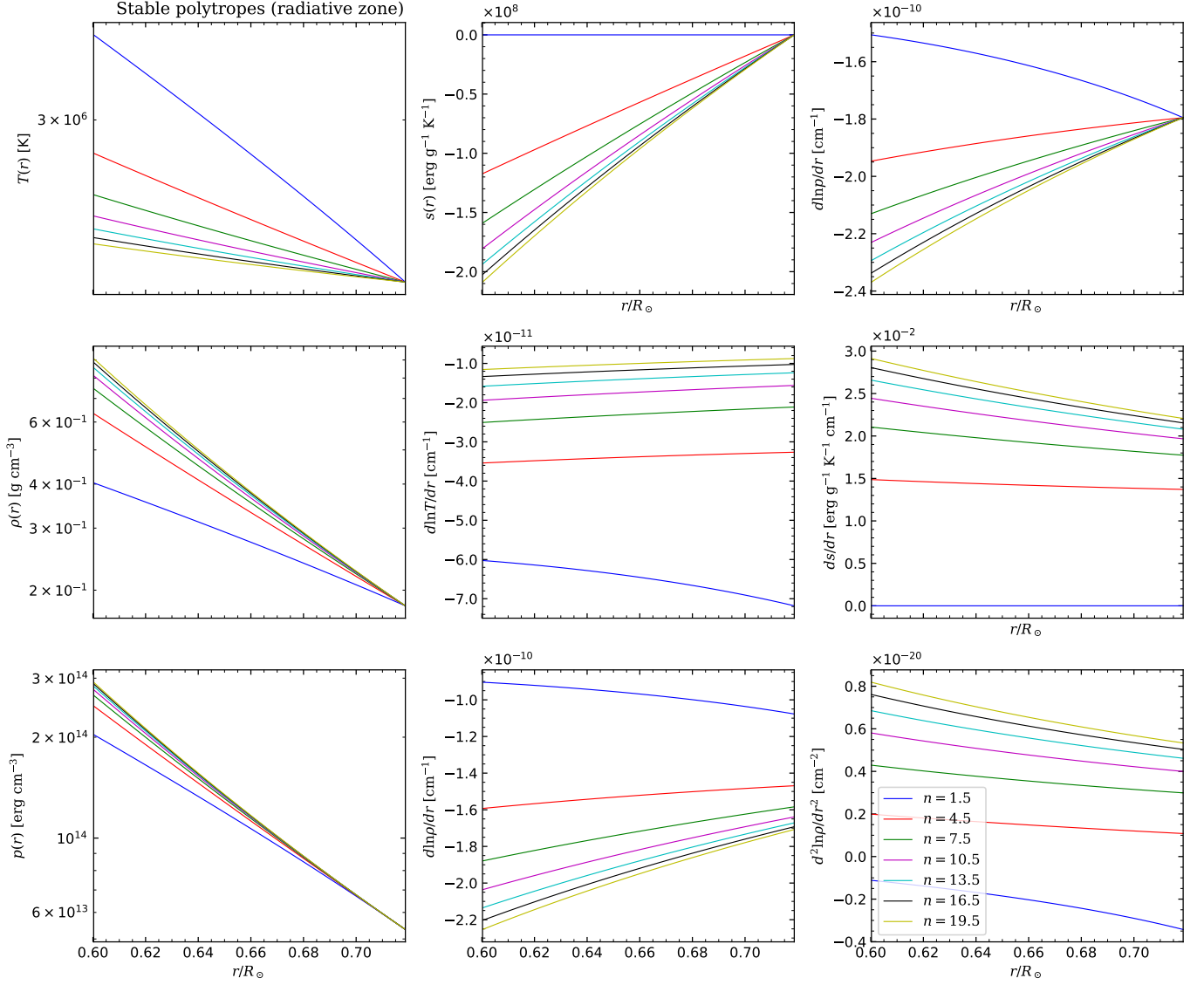


Figure 3: Sample of convectively stable polytropes for a solar-like RZ ($n \geq 1.5$). Similar to Figure 1, except the radius range is ($r_i = 4.176 \times 10^{10}$ cm, $r_m = 5.00 \times 10^{10}$ cm) instead of (r_m, R_o).

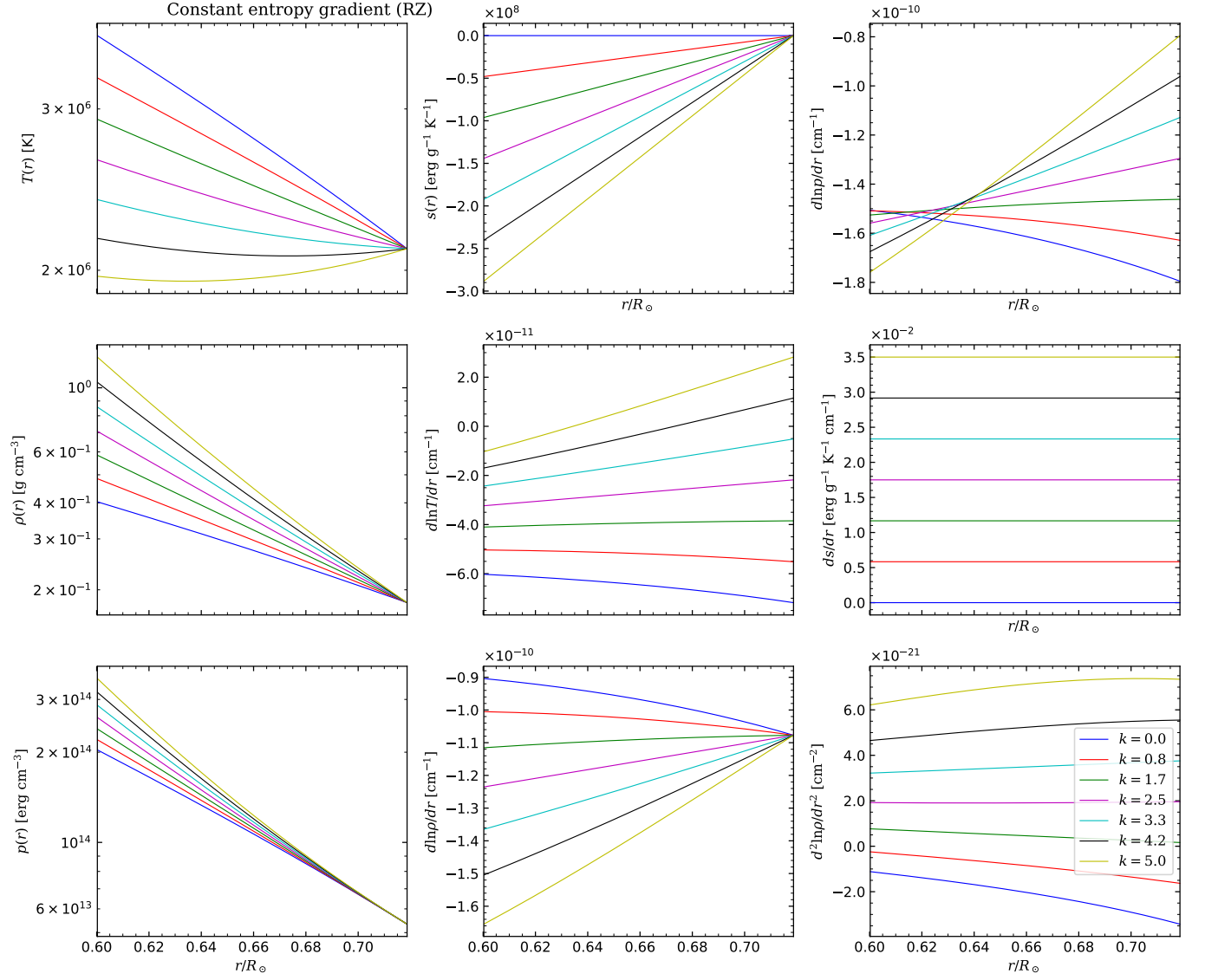


Figure 4: Sample of constant-entropy-gradient gas layers for a solar-like RZ, with values of k chosen to give entropy gradients more or less matching those of Figure 3.

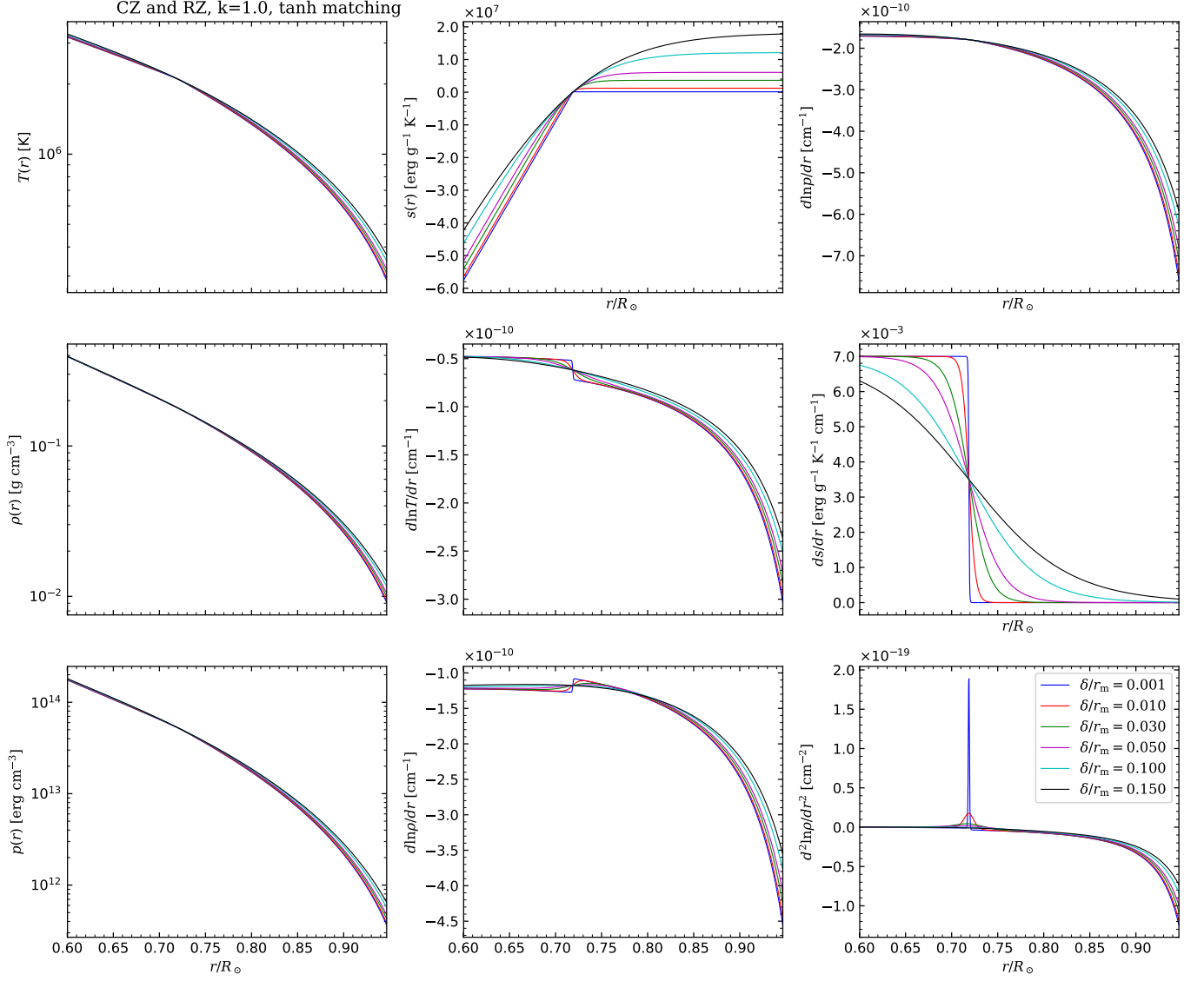


Figure 5: Combined CZ and RZ fluid layer with entropy gradient profiles matched via hyperbolic tangents—different transition thicknesses δ , with $k = 1$ in all cases.

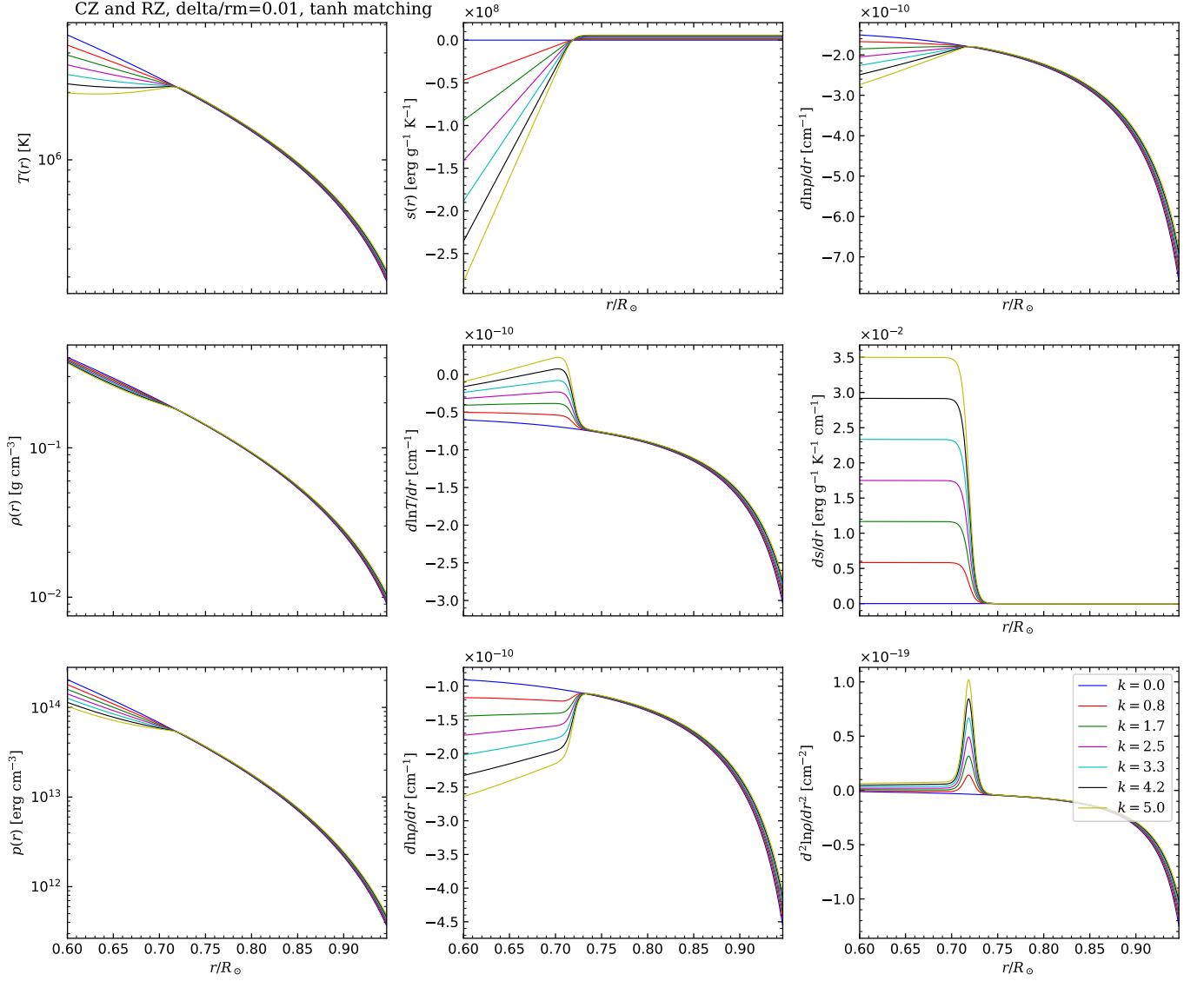


Figure 6: Combined CZ and RZ fluid layer with entropy gradient profiles matched via hyperbolic tangents—different levels of stiffness k , with $\delta/r_m = 0.01$ in all cases.

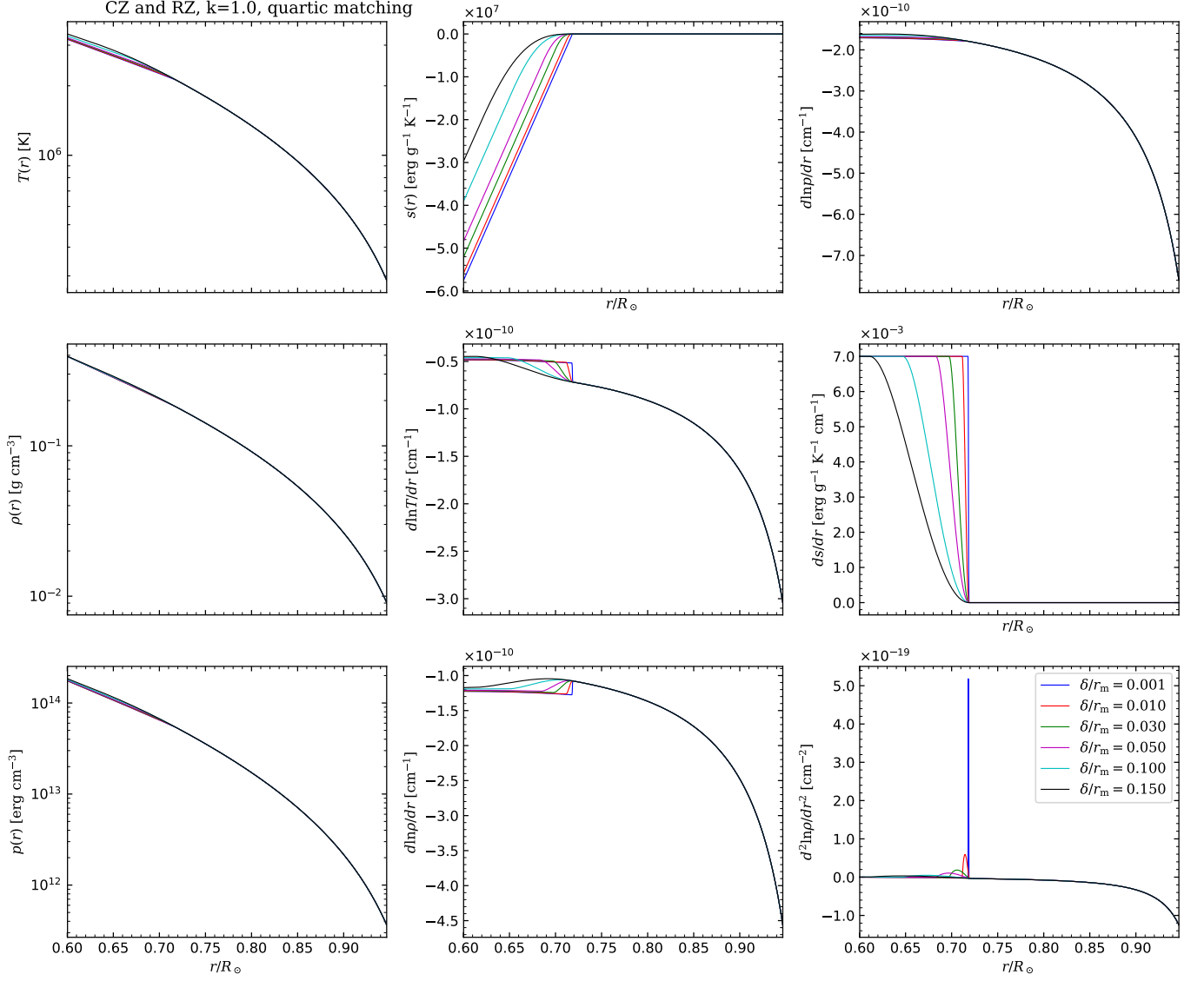


Figure 7: Combined CZ and RZ fluid layer with entropy gradient profiles matched via quartic functions—different transition thicknesses δ , with $k = 1$ in all cases.

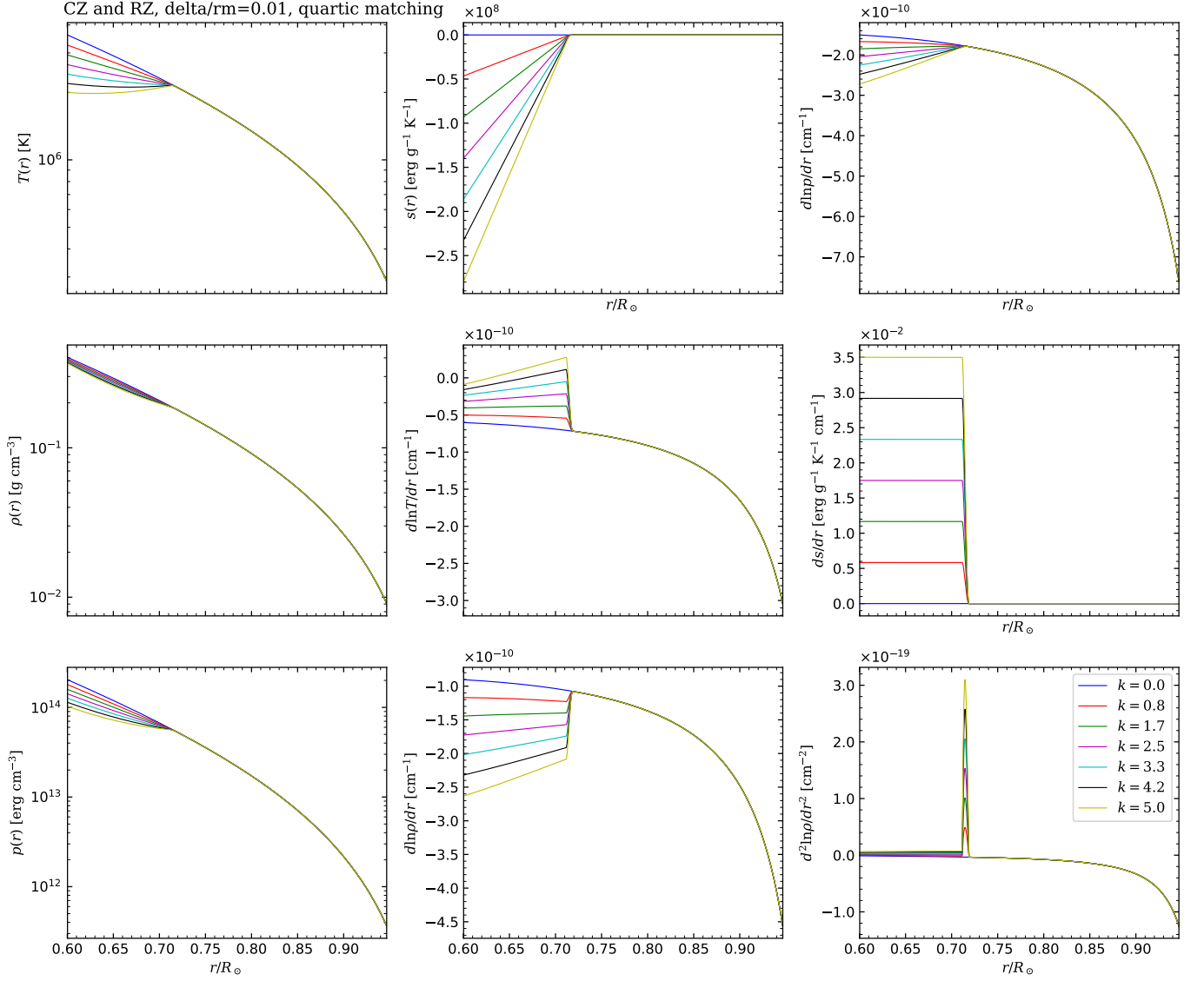


Figure 8: Combined CZ and RZ fluid layer with entropy gradient profiles matched via quartic functions—different levels of stiffness k , with $\delta/r_m = 0.01$ in all cases.