

Holographic dark energy models in $f(Q, T)$ gravity and cosmic constraint

Xuwei Zhang^{a,b}, Xiaofeng Yang^{a,b,*}, Yunliang Ren^{a,c}, Shuangnan Chen^{a,c} and Yangjun Shi^{a,d}

^aXinjiang Astronomical Observatory, Chinese Academy of Sciences, 150 Science 1-Street, Urumqi, 830011, China

^bSchool of Astronomy and Space Science, University of Chinese Academy of Sciences, No.19A Yuquan Road, Beijing, 100049, China

^cSchool of Physical Science and Technology, Xinjiang University, 666 Shenli Street, Urumqi, 830046, China

^dSchool of Physics and Astronomy, China West Normal University, 666 Shenli Street, Nanchong, 637002, China

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ABSTRACT

In this work, we investigate a modified gravity model $f(Q, T)$ with holographic dark energy. By considering the holographic principle as an extension to cosmology, we explore dark energy as a result of the entanglement entropy of the universe's horizon. We employ the Barrow holographic dark energy model, incorporating quantum corrections, and use the Hubble horizon for the infrared cutoff. Through parameter estimation using MCMC with the latest supernova, BAO, and Hubble data, we find that the model alleviates the Hubble tension and provides consistent results with the standard cosmological model. Additionally, the model successfully describes the accelerated expansion of the universe. Despite the complexity and additional parameters introduced by the model, it offers a promising framework for further exploration of dark energy within modified gravity theories.

1. Introduction

Over the past few decades, a series of groundbreaking discoveries in cosmology have significantly reshaped our understanding of the universe. In 1998, the accelerated expansion of the universe was first observed through studies of Type Ia supernovae [1, 2]. This discovery was later confirmed by a variety of other cosmological observations, including measurements of temperature anisotropies and polarization in the cosmic microwave background (CMB) radiation [3, 4], the peak length scale of baryon acoustic oscillations (BAO) [5, 6], the study of the large-scale structure (LSS) of the universe [7, 8], and direct measurements of the Hubble parameter using cosmic chronometers [9, 10]. These observations suggest the presence of a mysterious form of energy in the universe, known as dark energy (DE), which exhibits negative pressure and an increasing density. Dark energy acts like anti-gravity, yet its exact nature remains elusive.

Theoretical predictions and astronomical observations indicate that there may be a mysterious form of energy in the universe. This energy has the characteristics of negative pressure, and its density increases over time. This is considered to be the key factor driving the accelerated expansion of the universe, accounting for about three-quarters of the total energy of the universe [11, 12, 13]. In order to achieve such accelerated expansion, this form of energy needs to produce an anti-gravitational effect throughout the observable universe. However, ordinary baryonic matter neither has this equation of state nor can it explain such a large proportion of the cosmic energy component. Therefore, scientists have proposed and studied a variety of alternative theories and models to explore the nature of this cosmic acceleration phenomenon.

The simplest and most widely accepted theory is Λ CDM model, where Λ means cosmological constant predicted by Einstein [14]. Based on Λ CDM model, the latest observations suggest that our universe consists of 68.3% dark energy, 26.8% cold dark matter and 4.9% ordinary matter [4]. However, this model is not free from problems and the problems it is facing are cosmic coincidence, fine-tuning and the Hubble tension—a discrepancy between the value of the Hubble constant H_0 inferred from the CMB by the Planck satellite and that obtained from local measurements using Type Ia supernovae has sparked significant debate.

Another interesting attempt is to deviate from general relativity toward a modified form (detailed research progress can be reviewed in [15]). These theories assume that general relativity not work in large scale requiring a modification

✉ zhangxuwei@xao.ac.cn (X. Zhang); yangxiaofeng@xao.ac.cn (X. Yang)

ORCID(s): 0009-0002-9943-4062 (X. Zhang); 0000-0000-0000-0000 (X. Yang); 0000-0002-9659-4581 (Y. Ren); 0009-0001-5887-7652 (S. Chen); 0009-0007-8857-7371 (Y. Shi)

in action rather than standard Einstein-Hilbert action. The most well-known is $F(R)$ gravity which replaces the Ricci scalar R in the action by a general function $f(R)$ [16]. The $f(G)$ gravity theory is also a modified theory of gravity that introduces a correction to the GaussBonnet (GB) term G , allowing it to be arbitrary function $f(G)$ rather than remaining a constant [17, 18]. Another modified theory of gravity $f(T)$ extends the teleparallel equivalent of General Relativity (TEGR). It replaces the curvature scalar R in action with the torsion scalar T , derived from the Weitzenböck connection. Also shows some interpretations for the accelerating phases of our Universe [19, 20]. $f(Q)$ is generalized symmetric teleparallel gravity, with curvature and torsion both being zero, which is inspired by Weyl and Einstein's trial to unify electromagnetic and gravity. The geometric properties of gravity are described by "non-metricity". That is, the covariant derivative of the metric tensor is no longer zero (some detailed information can be found in review [21]). Harko et al. have proposed a new theory known as $f(R, T)$ gravity, where R stands for the Ricci scalar and T denotes the trace of energy-momentum tensor which presents a non-minimum coupling between geometry and matter [22]. Similar theories are introduced, $f(R, G)$ gravity proposed by Bamba et al. [23]; $f(Q, T)$ proposed by Xu et al. [24]; $f(Q, C)$ gravity [25]; $f(\mathcal{T}, T)$ proposed by [26]; $f(T, B)$ gravity [27, 28]; $f(R, T^2)$ proposed by Katirci et al. [29], etc.

Holographic dark energy is an famous alternative theory for the interpretation of dark energy, originating from the holographic principle proposed by 't Hooft [30]. Cohen et al. introduced the "UV-IR" relationship, highlighting that in effective quantum field theory, a system of size L has its entropy and energy constrained by the Bekenstein entropy bound and black hole mass, respectively. This implies that quantum field theory is limited to describing low-energy physics outside black holes [31]. After that, Li et al. proposed that the infrared cut-off relevant to the dark energy is the size of the event horizon and obtained the dark energy density can be described as $\rho_{de} = 3c^2 M_p^2 R_h^2$ where R_h is future horizon of our universe [32]. Although choose Hubble cut-off is a natural thought, but Hsu found it might lead to wrong state equation and be strongly disfavored by observational data [33].

Various attempts to reconstruct or discuss HDE in modified gravity have been completed by several authors. Wu and Zhu reconstructed HDE in $f(R)$ gravity [34]. Shaikh et al. discussed HDE in $f(G)$ gravity with Bianchi type 1 model [35]. Zubair et al. reconstructed Tsallis holographic dark energy models in modified $f(T, B)$ gravity [36]. Sharif et al. studied the cosmological evolution of HDE in $f(G, T)$ gravity [37] and Alam et al. investigated Renyi HDE in the same gravity [38]. Myrzakulov et al. reconstructed Barrow HDE in $f(Q, T)$ gravity [39]. Singh et al. and Devi et al. discussed HDE models respectively in $f(R, T)$ gravity and take cosmic constraint [40, 41].

In this article, we assume that our universe is described by $f(Q, T)$ gravity, with HDE as one component of the fluid. In Section 2, we briefly introduce non-Riemannian geometry and $f(Q, T)$ gravity. In Section 3, we incorporate HDE into the model and derive the solution. Section 4 presents the data and methods used to obtain constraints. In Section 5, we analyze the results and investigate the evolution of the models. Finally, Section 6 provides a brief conclusion.

2. $f(Q, T)$ gravity theory

Weyl in 1918 introduced an extension of Riemannian geometry, using a non-metricity tensor $Q_{\alpha\mu\nu} = \nabla_\alpha g_{\mu\nu} = -w_\alpha g_{\mu\nu}$, which describes how the length of a vector changes during parallel transport where w_α coincides with those of the electromagnetic potentials [42]. Weyl geometry can also be extended to so-called Weyl-Cartan geometry by considering the torsion of spacetime [24].

In Weyl-Cartan geometry, a connection can be decomposed into three independent parts: the Christoffel symbol $\hat{\Gamma}^\alpha_{\mu\nu}$, the contortion tensor $K^\alpha_{\mu\nu}$ and the disformation tensor $L^\alpha_{\mu\nu}$, so that the general affine connection can be expressed as [43]

$$\Gamma^\alpha_{\mu\nu} = \hat{\Gamma}^\alpha_{\mu\nu} + K^\alpha_{\mu\nu} + L^\alpha_{\mu\nu} \quad (1)$$

whereas

$$\hat{\Gamma}^\alpha_{\mu\nu} = \frac{1}{2} g^{\alpha\beta} (\partial_\mu g_{\beta\nu} + \partial_\nu g_{\beta\mu} - \partial_\beta g_{\mu\nu}) \quad (2)$$

$$K^\alpha_{\mu\nu} = \frac{1}{2} T^\alpha_{\mu\nu} + T_{(\mu}{}^\alpha{}_{\nu)} \quad (3)$$

$$L^\alpha_{\mu\nu} = \frac{1}{2} Q^\alpha_{\mu\nu} - Q_{(\mu}{}^\alpha{}_{\nu)} \quad (4)$$

are the standard Levi-civita connection of metric $g_{\mu\nu}$, contortion and disformation tensors respectively. In the above definitions, the torsion tensors and the non-metric tensor are introduced as follow

$$Q_{\rho\mu\nu} \equiv \nabla_\rho g_{\mu\nu} = \partial_\rho g_{\mu\nu} - \Gamma_{\rho\mu}^\beta g_{\beta\nu} - \Gamma_{\rho\nu}^\beta g_{\mu\beta} \quad (5)$$

$$T^\alpha_{\mu\nu} \equiv 2\Gamma^\alpha_{[\mu\nu]} = \Gamma^\alpha_{\mu\nu} - \Gamma^\alpha_{\nu\mu} \quad (6)$$

The non-metric tensor has two independent traces, namely $Q_\mu = Q_\mu^\alpha{}_\alpha$ and $\tilde{Q}^\mu = Q_\alpha^{\mu\alpha}$, so we can get quadratic non-metricity scalar as

$$Q = \frac{1}{4}Q_{\alpha\beta\mu}Q^{\alpha\beta\mu} - \frac{1}{2}Q_{\alpha\beta\mu}Q^{\beta\mu\alpha} - \frac{1}{4}Q_\alpha Q^\alpha + \frac{1}{2}Q_\alpha \tilde{Q}^\alpha \quad (7)$$

We consider the general form of the Einstein-Hilbert action for the $f(Q, T)$ gravity in the unit $8\pi G = 1$

$$S = \int \left(\frac{1}{2}f(Q, T) + \mathcal{L}_m \right) \sqrt{-g} d^4x \quad (8)$$

where f is an arbitrary function of the non-metricity, \mathcal{L}_m is known as matter Lagrangian, $g = \det(g_{\mu\nu})$ denotes determinant of metric tensor, and $T = g^{\mu\nu}T_{\mu\nu}$ is the trace of the matter-energy-momentum tensor, where $T_{\mu\nu}$ is defined as

$$T_{\mu\nu} = -\frac{2}{\sqrt{-g}} \frac{\delta(\sqrt{-g}\mathcal{L}_m)}{\delta g^{\mu\nu}} \quad (9)$$

Vary the action (8) with respect to the metric tensor $g_{\mu\nu}$ we can get

$$\delta S = \int \left(\frac{1}{2}\delta[f(Q, T)\sqrt{-g}] + \delta(\mathcal{L}_m\sqrt{-g}) \right) d^4x \quad (10)$$

$$= \int \frac{1}{2} \left(-\frac{1}{2}f g_{\mu\nu} \sqrt{-g} \delta g^{\mu\nu} + f_Q \sqrt{-g} \delta Q + f_T \sqrt{-g} \delta T - \frac{1}{2}T_{\mu\nu} \sqrt{-g} \delta g^{\mu\nu} \right) d^4x \quad (11)$$

In analogy to studies on torsionless $f(R)$ gravity and curvature-free $f(T)$ gravity, we can generalize the gravity to theories containing an arbitrary function of the non-metricity scalar i.e. $f(Q, T)$. Therefore, we consider the following action

$$-\frac{2}{\sqrt{-g}} \nabla_\alpha (f_Q \sqrt{-g} P^\alpha_{\mu\nu}) - \frac{1}{2} f g_{\mu\nu} + f_T (T_{\mu\nu} + \Theta_{\mu\nu}) - f_Q (P_{\mu\alpha\beta} Q^\alpha_{\nu\beta} - 2Q^{\alpha\beta}_{\mu} P_{\alpha\beta\nu}) = T_{\mu\nu} \quad (12)$$

Where tensor $\Theta_{\mu\nu}$ are defined as $g^{\alpha\beta} \delta T_{\alpha\beta} / \delta g^{\mu\nu}$ and $P^\alpha_{\mu\nu}$ is the super-potential of the model (detailed discussion found in [24]). In the case of a globally vanishing affine connections, the non-metricity tensor depends on the metric only and Einstein's GR action is recovered. This occurs under the choice of the coincidence gauge, in which the origin of spacetime and that of the tangent space coincide. In the coincident gauge with $\Gamma^\alpha_{\mu\nu} = 0$ we have $Q = 6H^2$ (detailed discussion can be found in [44]).

Assuming that the Universe is described by the isotropic, homogeneous and spatially flat Friedmann-Lemaitre-Robertson-Walker (FLRW) spacetime, given by as line element

$$ds^2 = -N^2(t)dt^2 + a^2(t)\delta_{ij}dx^i dx^j \quad (13)$$

where $a(t)$ is the cosmic scale factor used to define the Hubble expansion rate $H = \dot{a}/a$ and the lapse function $N(t)$ used to define dilation rates $\tilde{T} = \dot{N}/N$ (for standard case $N(t) = 1$). To derive Friedmann equations describing the cosmological evolution, we assume that the matter content of the Universe consists of a perfect fluid, whose energy-momentum tensor is given by $T^\mu_\nu = \text{diag}(-\rho, p, p, p)$ and tensor Θ^μ_ν is expressed as $\text{diag}(2\rho + p, -p, -p, -p)$. Then use the line element (13) and field equation (12), we can get Friedmann equations

$$\rho = \frac{f}{2} - 6f_Q H^2 - \frac{2f_T}{1+f_T} (\dot{f}_Q H + f_Q \dot{H}) \quad (14)$$

$$p = -\frac{f}{2} + 6f_Q H^2 + 2(\dot{f}_Q H + f_Q \dot{H}) \quad (15)$$

where $f(Q, T)$ is simplified to f , and $f_Q = \partial f / \partial Q$, $f_T = \partial f / \partial T$, $\dot{f}_Q = \partial f_Q / \partial t$. By the usage of Eq. (7) and the line element (13), there exists following relationship (The detailed derivation can be found in the appendix of [24])

$$Q = 6H(t)^2 \quad (16)$$

The equation of state (EoS) parameter is given by

$$w = \frac{p}{\rho} = -1 + \frac{4f_Q H + f_Q \dot{H}}{(1 + f_T)(f - 12f_Q H^2) - 4f_T(\dot{f}_Q H + f_Q \dot{H})}, \quad (17)$$

where ρ and p denote the total energy density and pressure of the universe. Since we mainly focus in the late universe, the contribution of radiation can be ignored, so we only care about baryonic matter and holographic dark energy fluid.

$$\rho = \rho_m + \rho_{de}, \quad p = p_m + p_{de} \quad (18)$$

In our universe, the condition $w < -1/3$ ensures accelerated expansion.

Effective EoS parameter denote geometry qualities.

$$\rho_{\text{eff}} = 3H^2 = \frac{f}{4f_Q} - \frac{1}{2f_Q}[(1 + f_T)\rho + f_T p] \quad (19)$$

$$-p_{\text{eff}} = 2\dot{H} + 3H^2 = \frac{f}{4f_Q} - \frac{2\dot{f}_Q H}{f_Q} + \frac{1}{2f_Q}[(1 + f_T)\rho + (2 + f_T)p] \quad (20)$$

The effective energy density ρ_{eff} and pressure p_{eff} described above highlight the coupling between geometry and matter within the $f(Q, T)$ gravity framework. The presence of f_T explicitly links the geometric modifications to the energy density ρ and pressure p of the matter content. Additionally, the term involving \dot{f}_Q introduces time dependence in the coupling, suggesting a dynamical interplay between the evolution of spacetime geometry and matter distribution.

Furthermore, the effective EoS using Eq.(19)(20) can be written as

$$w_{\text{eff}} = \frac{p_{\text{eff}}}{\rho_{\text{eff}}} = \frac{-2\dot{H} - 3H^2}{3H^2} = -\frac{f - 8\dot{f}_Q H + 2[(1 + f_T)\rho + (2 + f_T)p]}{f - 2[(1 + f_T)\rho + f_T p]} \quad (21)$$

3. Cosmic solutions with holographic dark energy

The holographic principle sets an upper limit on the entropy of the universe. In the HDE model, the energy density of dark energy is typically expressed as ([32])

$$\rho_{de} = 3c^2 M_p^2 L^{-2} \quad (22)$$

where L is the characteristic length scale of the universe, and c is free parameter, M_p denotes planck mass here we set it as 1. The Hubble horizon is considered the simplest option. In addition, the particle horizon L_p or the future event horizon L_F are also considered reasonable options. Consider a simple case, the HDE energy density in Hubble cut-off can be described as

$$\rho_{de} = 3c^2 H(t)^2 \quad (23)$$

Another HDE model called barrow holographic dark energy (BHDE) generalizes holographic entropy that arises from quantum-gravitational effects which deform the black-hole surface by giving it an intricate, fractal form . In this case, the HDE energy density can be define as

$$\rho_{de} = 3c^2 H(t)^{2-\Delta} \quad (24)$$

here a new exponent Δ quantifies the quantum-gravitational deformation, with $\Delta = 0$ coming back to the standard Bekenstein-Hawking entropy, and with $\Delta = 1$ corresponding to the most intricate and fractal structure [45].

In order to incorporate holographic dark energy in the modified gravitational universe, we consider a simple form of f

$$f(Q, T) = mQ^n + \alpha T \quad (25)$$

where $Q = 6H^2$, $T = -\rho + 3p$, m , n and α are constants. So that we can derive $f_Q = \alpha n Q^{n-1} = \alpha n 6^{n-1} H^{2n-2}$, $f_T = \beta$, $\dot{f}_Q = 2\alpha n(n-1)6^{n-1} H^{2n-3} \dot{H}$.

We also introduce the deceleration factor, which describes the acceleration or deceleration of the late universe depending upon its value, is defined as

$$q = \frac{d}{dt} \frac{1}{H} - 1 = -\frac{\ddot{a}}{\dot{a}^2} = -\frac{\dot{H}}{H^2} - 1 = (1+z) \frac{1}{H(z)} \frac{dH(z)}{dz} - 1 \quad (26)$$

In order to understand the characteristic properties of the dark energy and more easier to get a analytical solution, we need to parameterize EoS of HDE as a constant parameter

$$w_{de} = \frac{p_{de}}{\rho_{de}} \quad (27)$$

Use Eq.(19) and (20) we can solve the energy density of fluid component in field equation

$$\rho_{de} = \frac{m(2n-1)(6H(t)^2)^{n-1}((3\alpha+2)nH'(t) + 3(\alpha+1)H(t)^2)}{(2\alpha^2 + 3\alpha + 1)w_{de}} \quad (28)$$

$$\rho_m = \frac{m(2n-1)(6H(t)^2)^{n-1}(n(\alpha(w_{de}-3) - 2)H'(t) - 3(\alpha+1)(w_{de}+1)H(t)^2)}{(2\alpha^2 + 3\alpha + 1)w_{de}} \quad (29)$$

This is a second order differential equation and it depends on t . In order to get cosmological solution, there is also a simple relation between $H(t)$ and $H(z)$

$$\dot{H}(t) = \frac{d}{dt} H(t) = -\frac{dH(z)}{dz} H(z)(1+z) \quad (30)$$

Combine equation (28) (29) and (30) and relation (30), we can get the equation follows

$$3c^2 H(z)^2 = \frac{m6^{n-1}(2n-1)(H(z)^2)^{n-1}(3(\alpha+1)H(z)^2 - (3\alpha+2)n(z+1)H(z)H'(z))}{(2\alpha^2 + 3\alpha + 1)w_{de}} \quad (31)$$

In principle, we can get the form of $H(z)$ through the solution of above differential equation. However, solving higher-order differential equations analytically is difficult. So we consider $n = 1$ firstly to simplify calculation and set the initial condition $H(z=0) = H_0$ which denotes value of the Hubble parameter at present, get analytically solution as follow

$$H(z) = H_0(1+z)^{\frac{3(1+\alpha)(m-c^2w_{de}(1+2\alpha))}{m(2+3\alpha)}} \quad (32)$$

We thus obtain the power-law evolution of the Universe which avoids the big-bang singularity similar to the $f(R, T)$ situation in ([46]). In this case, the deceleration factor is

$$q = -1 + \frac{3(\alpha+1)(m - (2\alpha+1)c^2w_{de})}{(3\alpha+2)m} \quad (33)$$

Here we have a constant deceleration factor depending on the parameter. When we select a particular parameter value, it will show acceleration or deceleration characteristics. If $q < 0$, it will show the characteristics of accelerated expansion.

However, since the deceleration factor is time independent, there is no phase transition in such a universe. So we can use Eq. (24) to obtain a tighter UV limit, if we set $\Delta = 1$ we can get

$$H(z) = H_0(1+z)^{\frac{3(\alpha+1)}{3\alpha+2}} - (1+2\alpha)c^2w_{de} \left((1+z)^{\frac{3(\alpha+1)}{3\alpha+2}} - 1 \right) \frac{1}{m} \quad (34)$$

In this case, the deceleration factor is

$$q = \frac{(2\alpha+1)c^2w_{de} \left(3\alpha + (z+1)^{\frac{3(\alpha+1)}{3\alpha+2}} + 2 \right) - H_0m(z+1)^{\frac{3(\alpha+1)}{3\alpha+2}}}{(3\alpha+2) \left((2\alpha+1)c^2w_{de} \left((z+1)^{\frac{3(\alpha+1)}{3\alpha+2}} - 1 \right) - H_0m(z+1)^{\frac{3(\alpha+1)}{3\alpha+2}} \right)} \quad (35)$$

in other situation, if $n \neq 1$ or $\Delta \neq 1$, higher-order differential equations are difficult to solve analytically, so we can only obtain numerical solutions through complex machine computing. For comparison, we also consider the case where $\alpha = 1$ that reduces it to the minimum matter coupling.

4. Observational data and methodology

In this work, we estimate the cosmological parameters of the model by employing a Markov Chain Monte Carlo (MCMC) method based on the minimization of the chi-square function, χ^2 [47]. The chi-square function is given by:

$$\chi^2 = \sum_i \left(\frac{D_i - T_i(\theta)}{\sigma_i} \right)^2, \quad (36)$$

where D_i represents the i -th data point, $T_i(\theta)$ is the theoretical prediction for the corresponding quantity, and σ_i is the error associated with the i -th data point. Here, θ denotes the vector of model parameters.

For our analysis, we combine three independent observational datasets:

1. Baryon Acoustic Oscillations (BAO): The BAO measurements provide a standard ruler for distance measurements in the universe. We use the data from the SDSS Baryon Oscillation Spectroscopic Survey (BOSS), Dark Energy Spectroscopic Instrument (DESI) first year data and 6dF Galaxy Survey (6dFGS) to constrain the cosmological parameters. The comoving horizon distance, the transverse comoving distance and the volume-averaged distance combining line-of-sight and transverse distances defined as follow

$$D_H = \frac{c}{H(z)} \quad (37)$$

$$D_M = \frac{D_L}{1+z} \quad (38)$$

$$D_V = \left[\frac{cz}{H(z)} \right]^{1/3} \left[\frac{D_L}{1+z} \right]^{2/3} \quad (39)$$

Where D_L is the luminosity distance. When scaled by the sound horizon at the drag epoch r_d , ratios such as D_H/r_d , D_M/r_d , and D_V/r_d serve as important observables for constraining cosmological models and testing the standard model of cosmology.

2. Chronometers Data (OHD): The Hubble parameter measurements, known as the chronometers data, provide independent estimates of the Hubble parameter $H(z)$ at various redshifts. These data serve as an important probe of the expansion rate of the universe. We choose the dataset from [52]

3. Type Ia Supernovae (SNIa) Data: Type Ia supernovae (SNIa) are considered standard candles because When the light curve reaches its maximum, the absolute luminosity is almost the same. The distance modulus μ can be obtained according to the following formula

$$\mu_{obs} = m - M \quad (40)$$

On the other hand, we can get the theoretical distance modulus from the cosmological model

$$\mu_{th}(z) = 5 \log_{10} d_L(z) + 25 + M_b \quad (41)$$

| Survey | z_{eff} | D_M/r_d | D_H/r_d | D_V/r_d |
|------------------------------------|------------------|------------------|------------------|------------------|
| 6dFGS | 0.106 | | | 2.98 ± 0.13 |
| SDSS MGS | 0.15 | | | 4.51 ± 0.14 |
| SDSS DR12 | 0.38 | 10.27 ± 0.15 | 24.89 ± 0.58 | |
| SDSS DR12 | 0.51 | 13.38 ± 0.18 | 22.43 ± 0.48 | |
| SDSS DR16 LRG | 0.70 | 17.65 ± 0.30 | 19.78 ± 0.46 | |
| SDSS DR16 ELG | 0.85 | 19.50 ± 1.00 | 19.60 ± 2.10 | |
| SDSS DR16 QSO | 1.48 | 30.21 ± 0.79 | 13.23 ± 0.47 | |
| SDSS DR16 Ly α -Ly α | 2.33 | 37.60 ± 1.90 | 8.93 ± 0.28 | |
| SDSS DR16 Ly α -QSO | 2.33 | 37.30 ± 1.70 | 9.08 ± 0.34 | |
| DESI BGS | 0.30 | | | 7.93 ± 0.15 |
| DESI LRG1 | 0.51 | 13.62 ± 0.25 | 20.98 ± 0.61 | |
| DESI LRG2 | 0.71 | 16.85 ± 0.32 | 20.08 ± 0.60 | |
| DESI LRG+ELG | 0.93 | 21.71 ± 0.28 | 17.88 ± 0.35 | |
| DESI ELG | 1.32 | 27.79 ± 0.69 | 13.82 ± 0.42 | |
| DESI QSO | 1.49 | | | 26.07 ± 0.67 |
| DESI Ly α -QSO | 2.33 | 39.71 ± 0.94 | 8.52 ± 0.17 | |

Table 1

BAO dataset. This table is referenced from [48], including 6dFGS data [49], SDSS data [50]) and DESI2024 BAO data [51]

where M_b denotes the absolute luminosity of SNIa and the luminosity distance is defined as

$$d_L(z) = \frac{c}{H_0} (1+z) \int_0^z \frac{dz'}{E(z')} \quad (42)$$

In this paper, we use Pantheon+ dataset who comprises 1701 SNIa samples, an increase from the 1048 samples in Pantheon dataset. Pantheon+ dataset consists of 1701 light curves of 1550 spectroscopically confirmed SNIa within the redshift range of $0.001 < z < 2.26$ [53, 54].

The combined likelihood function \mathcal{L} is then constructed by multiplying the individual likelihoods of each dataset:

$$\mathcal{L} = \mathcal{L}_{\text{BAO}} \times \mathcal{L}_{\text{OHD}} \times \mathcal{L}_{\text{SNIa}} \quad (43)$$

it is implied that

$$\chi_{\text{tot}}^2 = \chi_{\text{BAO}}^2 + \chi_{\text{OHD}}^2 + \chi_{\text{SNIa}}^2 \quad (44)$$

In order to test the statistical significance of our constraints we implement the AIC criterion. To complete our statistical study, we also use the selection method named the Bayesian Information Criterion (BIC). To calculate the value of BIC for each model we use

5. Results and analysis

The results of the parameter constraints obtained in our analysis are summarized in Table 2. We find the following 95% confidence limits for the cosmological and model parameters: the Hubble constant is $H_0 = 67.7^{+2.5}_{-2.7}$ km/s/Mpc, which is consistent with recent Planck measurements, though slightly lower. The parameter c , which governs the modification to the gravitational dynamics, is constrained to $c = -8.8^{+1.9}_{-2.0}$, indicating a significant deviation from the standard cosmological model. The parameter a , related to the evolution of the modified equation of state, is found to be $a = -0.180^{+0.065}_{-0.060}$, suggesting a non-negligible deviation from a simple cosmological constant. The inverse mass parameter $1/m$ is constrained to $1/m = -1.13^{+0.38}_{-0.46}$, reflecting the sensitivity of the model to the mass scale of new physics. The equation of state parameter for dark energy, w_{de} , is constrained to $w_{\text{de}} = -0.79^{+0.25}_{-0.28}$, indicating that dark energy is close to but slightly less than the value for a cosmological constant. The absolute magnitude of the reference galaxy, M_b , is $M_b = -19.397^{+0.079}_{-0.086}$, with a narrow error range consistent with the expected value for the sample of galaxies considered. Finally, the sound horizon at the drag epoch, r_d , is measured to be $r_d = 145.2^{+5.7}_{-5.2}$ Mpc, in

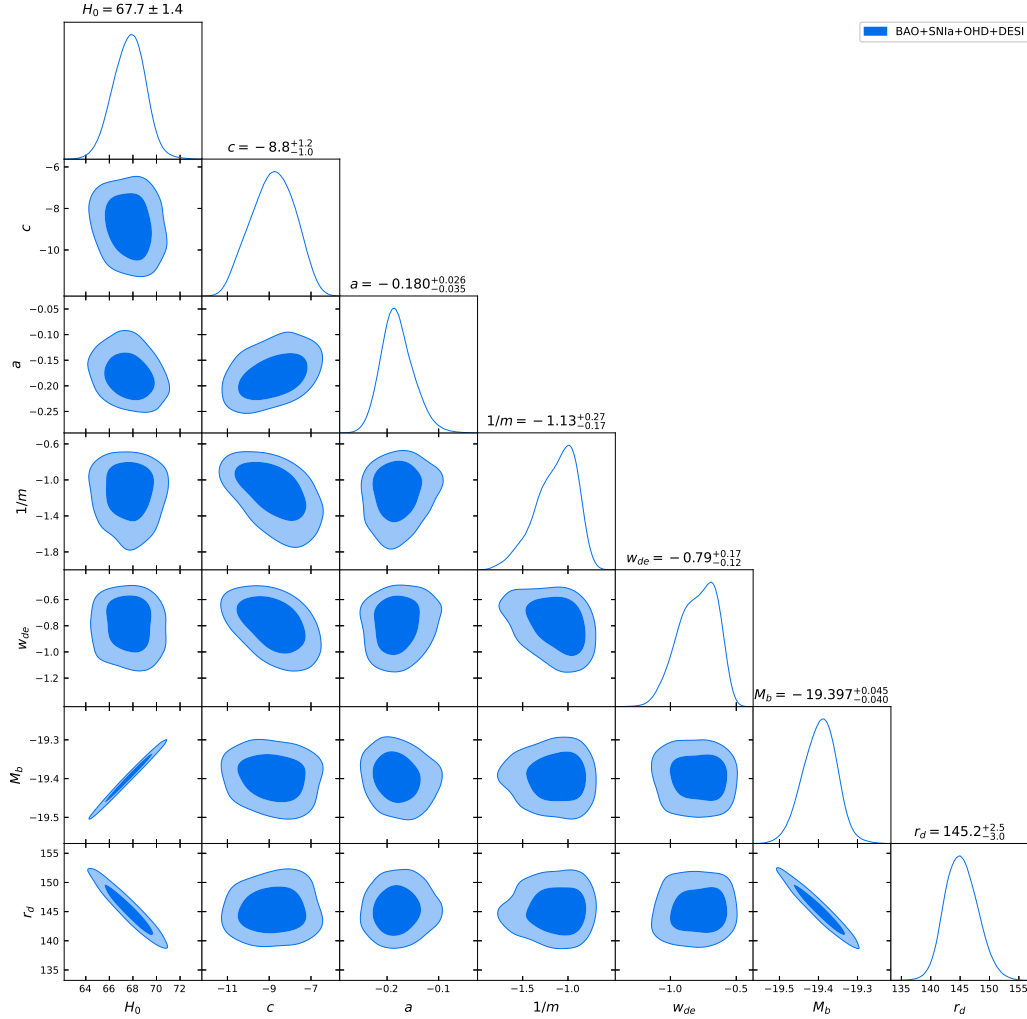


Figure 1: The 1σ and 2σ confidence contours and the 1D posterior distributions obtained from MCMC constraint of HDE in $f(Q, T)$ gravity using BAO+SNla+OHD+DESI data

agreement with current measurements from the cosmic microwave background (CMB) and other large-scale structure surveys. These results provide a comprehensive view of the cosmological parameters within the modified gravity framework, indicating a viable model that is consistent with observational data while deviating from the standard cosmological model, particularly in the modification to the dark energy equation of state and gravitational dynamics. The results show that our results may alleviate the Hubble tension.

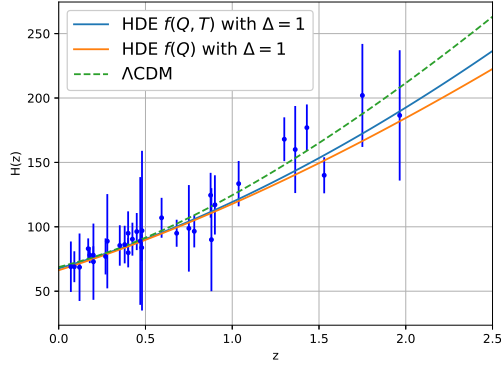
6. Conclusion

In this paper, we discuss the evolution of holographic dark energy in the context of a non-metric modified gravity theory $f(Q, T)$. While this may seem like an overcomplicated assumption, in reality, if we live in a non-metric space

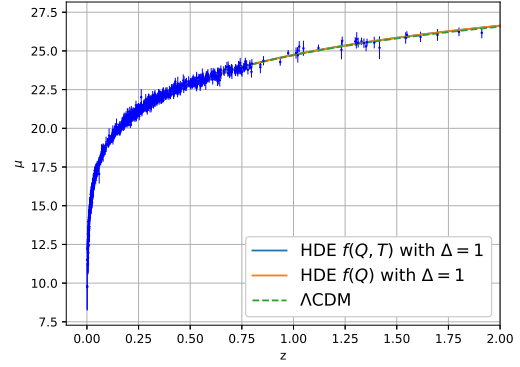
Table 2

Results of Constraint on parameters with prior ranges and 95% credible limits

| Parameter | Prior | 95% limits |
|-----------|------------|-----------------------------|
| H_0 | [50, 100] | $67.7^{+2.5}_{-2.7}$ |
| c | [-20, 0] | $-8.8^{+1.9}_{-2.0}$ |
| a | [-1, 0] | $-0.180^{+0.065}_{-0.060}$ |
| $1/m$ | [-2, 0] | $-1.13^{+0.38}_{-0.46}$ |
| w_{de} | [-1.5, 0] | $-0.79^{+0.25}_{-0.28}$ |
| M_b | [20, 30] | $-19.397^{+0.079}_{-0.086}$ |
| r_d | [100, 200] | $145.2^{+5.7}_{-5.2}$ |



(a) name1



(b) Data points of supernova distance modulus and redshift and the fit of each model

Figure 2: Main name

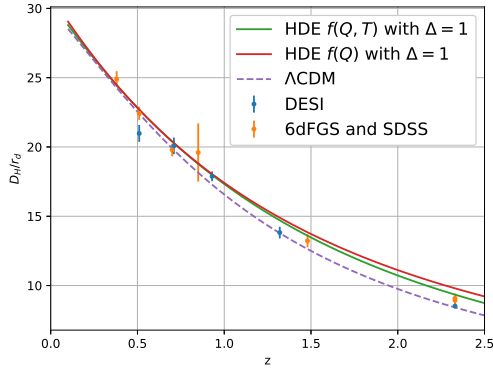
(such as Weyl-Cartan geometry) rather than a Riemannian space, we cannot simply assume that dark energy no longer exists as a fluid. Instead, we continue to treat it as one of the driving forces behind cosmic acceleration, even in such a complex universe.

We consider holographic dark energy due to its solid theoretical foundation and its interpretation as an extension of the holographic principle in cosmology. It provides a compelling explanation for the origin of dark energy, which arises from the entanglement entropy of the cosmic horizon. To refine our model, we explore a generalized version of holographic dark energy, known as Barrow holographic dark energy, which may better capture quantum corrections. For the infrared cutoff, we choose the Hubble horizon, as it is the simplest and most natural choice, although other horizons could also be considered.

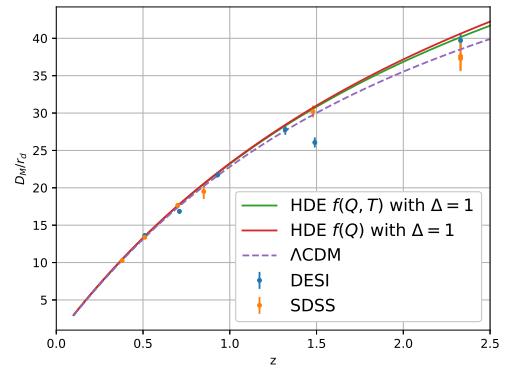
For analytical simplicity, we focus on a few toy models that allow us to explore certain special cases: $\Delta = 0$ for the loosest constraints and $\Delta = 1$ for the tightest limits on energy density, while discussing the potential effects of non-minimal matter coupling.

To validate the model, we performed parameter estimation using recent supernova data, BAO data, and direct measurements of the Hubble parameter. By employing the MCMC method, we obtained estimates for the model parameters. Our results show that the model can effectively alleviate the Hubble constant tension. Specifically, we find that the Hubble constant is ..., in agreement with the standard cosmological model's result of We also study the evolution of the universe under this model and observe that the deceleration factor and the effective equation of state parameter indicate accelerated expansion, consistent with current observations.

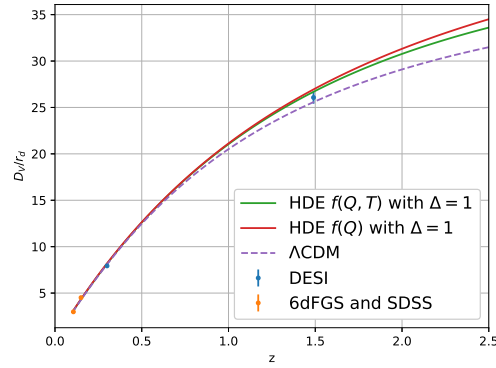
Although this model is not yet conclusive due to the introduction of numerous parameters and significant assumptions, it provides an interesting framework for thinking about cosmic evolution. The results and analysis presented here offer valuable insights, and further exploration of more complex scenarios will be necessary in future studies.



(a) name1



(b) name2

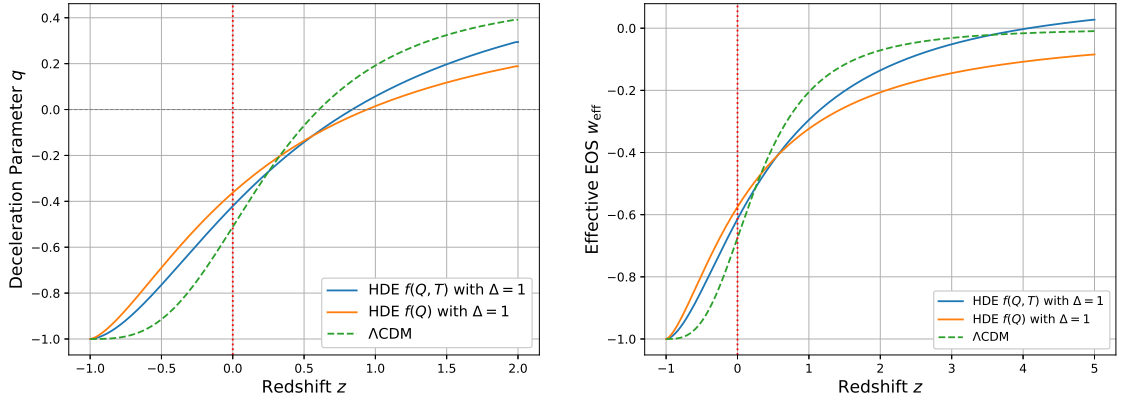


(c) name3

Figure 3: Main name

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(a) The evolution of the models' deceleration factor which reflects the accelerating expansion phase in $z \approx 0.8$

(b) The effective EOS, accounting for all components, provides a direct reflection of the cosmic expansion history and ultimately approaches -1.

Figure 4: Evolution of deceleration factor and effective EOS

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