

Practical 1

Aurora Borealis

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1 Objectives

Polar lights, also known as Aurora Polaris, is the collective term for the aurora seen over the northern and southern poles of the Earth. They are referred to as northern lights (Aurora Borealis) and southern lights (Aurora Australis). Research of the polar lights has been carried out for many decades; there are many theories about how this spectacular display is created. [1]

The primary objective of this practical was to understand the phenomena of aurora and be able to classify it according to different criteria. As a part of this assignment, it was also required to study and correlate the riometer and magnetometer data during the occurrence of auroras.

2 Classification

TBD The images in Figure 1 and Figure 2 were taken from the IRF website. The aurora in the Figure 1 has been classified as a_4fR_2B2a because of the following reasons.

- **Condition:** It has been classified as active a_4 because the aurora appeared as a whole and new form appeared after the previous one had faded away.
- **Qualifying Symbol:** It is fragmented in the upper left corner and hence it is fragmentary f
- **Structure:** The auroras has form which appears as rays aligned to the magnetic field. It also appears as a gigantic curtain or drapery. Hence it is classified as of type R_2
- **Brightness and Colour:** The brightness index seems to be of type B2 and since the colour seen are green with patches of yellow and purple it is classified to be of type a



Figure 1: Image taken from IRF Website

The aurora in the Figure 2 has been classified as $qHP3c$ because of the following reasons.

- **Condition:** The aurora in Figure 2 appears to be quiet and hence it is of type q

- **Structure:** The aurora has no specific internal structure and has a uniform brightness and colour change only slightly from one side to the other. Hence it is of homogeneous H type.
- **Brightness and Colour:** The light intensity is comparable with brightly moonlit cirrus clouds or cumulus cloud and so it is of type P3. Since the colour is green it is of type c.



Figure 2: Second Image taken from IRF Website

3 Auroral Forecast

The purpose of this part of the assignment is to analyse data from different sensors and predict whether it is suitable conditions for seeing Aurora Borealis or not. However, by the time this report has been written (April 26, 2012) it was not possible to see the northern lights because it didn't get dark enough during the night time. For this reason, this section contains the analysis of the data taken from previous month and is compared to the images from the Auroral Large Imaging System (ALIS).

The data for analysis was taken from May 28, 2012 because it was quite a spectacular Aurora that day. As a source of images and scientific data the website of Swedish Institute of Space Physics (IRF) was used. The scientific data implies magnetometer and riometer measurements.

3.1 Magnetometer

Magnetometer is often used to measure long- and short-term variations. The Kiruna flux-gate magnetometers are used to measure short-term variations, which are due to electric currents flowing in the ionosphere and induced in the surface layers in the solid Earth. The time resolution for the data measurement of these particular sensors is 20 seconds. The key properties of the magnetic disturbance level are:

- K-indices: Quasi-logarithmic scale from 0 (very quiet) through 3 (moderately disturbed) to 9 (very disturbed).

- A_p index: A linear, daily index ranging from 0 (very quiet) through 15 (moderately disturbed) to 400 (very disturbed).
- A_E index: Index of auroral activity, derived from horizontal component of magnetic field variation at 10-12 auroral-zone observatories.
- D_{st} index: Index of ring-current (and global magnetic storms), derived from near equator observatories.

The flux-gate data from Kiruna's magnetometers is calibrated against the absolute measurements each month. Afterwards, quiet days are identified and the K-indices are calculated. In the geomagnetic measurements disturbance variations might be a sign of an aurora.

3.2 Riometer

The Relative Ionosphere Opacity meters (riometers) are used to measure the absorption of radio-noise from the starts in the ionosphere and they operate at 25-50 Mhz frequency. Radio waves of this frequency are absorbed at the altitude of 60-110 km if there is a significant number of electrons present in ionosphere. Riometers perform continuous observations of the noise-level and observe decreases in noise relative to the quiet-day level when there is increased ionization due to the precipitation of the high energy electrons and protons. The typical auroral signature on the riogram is a sharp dip of few dB, followed by the slow regrowth. IRF has two riometers working at 30 and 38 MHz.

3.3 Analysis

The measurements taken from the archive for 28th of March are shown in the figures below, where Figure 3 represents magnetometer measurements and Figure 4 represents riogram.

As it is seen in the magnetogram, there were quite some fluctuations from 00:00 to 08:00 at night of 28th of March. However, images are not recorded though the entire night because at the certain point it becomes too bright to see the aurora. Riometer has also some fluctuations during that time. Therefore, we can assume that there was a high probability of auroral display. However, to make sure that our assumptions are correct we decided to focus on the disturbances, occurred in the period from 20:00 to 22:00. The magnetometer clearly shows the variations in the magnetic field and the riometer has a dip, followed by the ramp of magnetic field recovery. Based on these observations, we conclude that there should have been a strong aurora during that period.

To check if the predictions were correct or not the pictures from the archive were taken for the period we are interested in. Compilation of 12 images from ALIS is shown in Figure 5. According to the datagrams the auroral peak was around 21:00. On the ALIS images, the one taken at 20:51 is the brightest which is approximately at the same time as expected. To sum up, it is seen that the measurements reflect the real auroral probability correctly.

4 Results and Conclusions

TBD for the first part

The date of the experiment was chosen based on personal observations. During the experiment, datagrams from the magnetometer and riometer were compared to the images taken by ALIS. All the data was fetched from the IRF archives. The analysis showed that the highest probability of the aurora took place at the same time as there was the brightest aurora according to the images. As a result, we conclude that we have understood the principle how the predictions are made based on the data analysis and what processes are responsible for studied behavior.

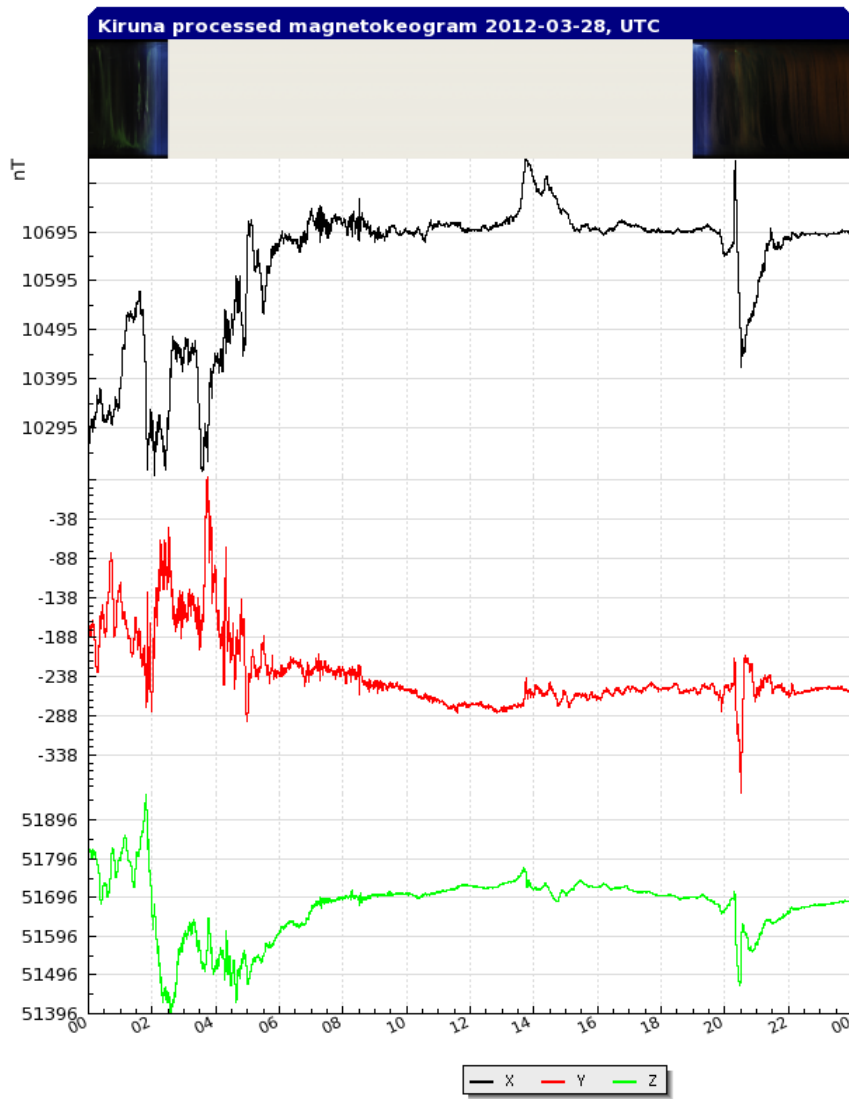


Figure 3: Magnetogram for the 28th of March, 2012.

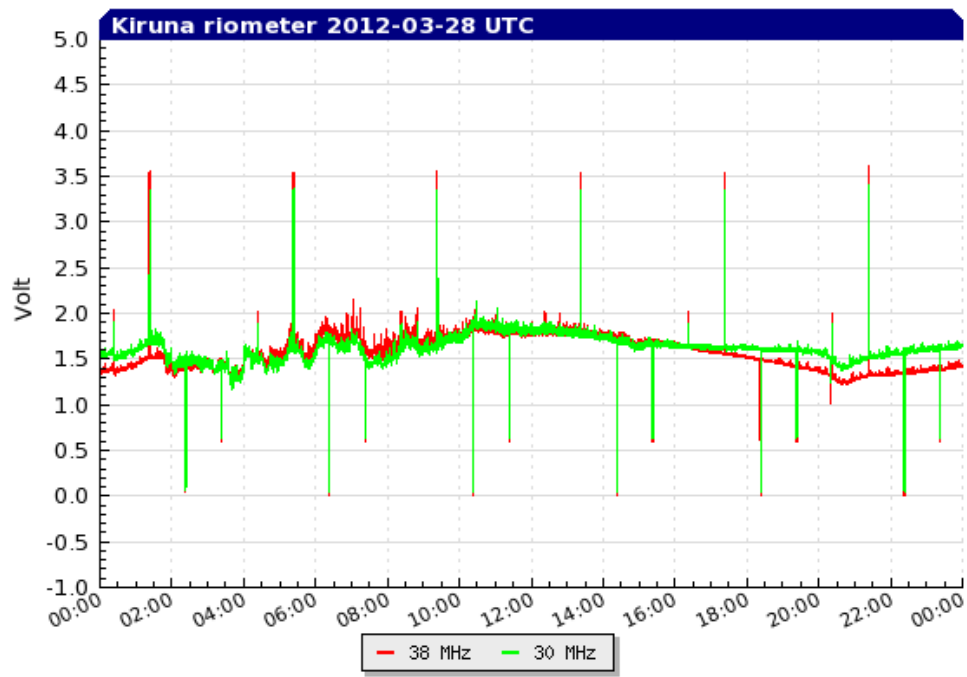


Figure 4: Riogram for the 28th of March, 2012.

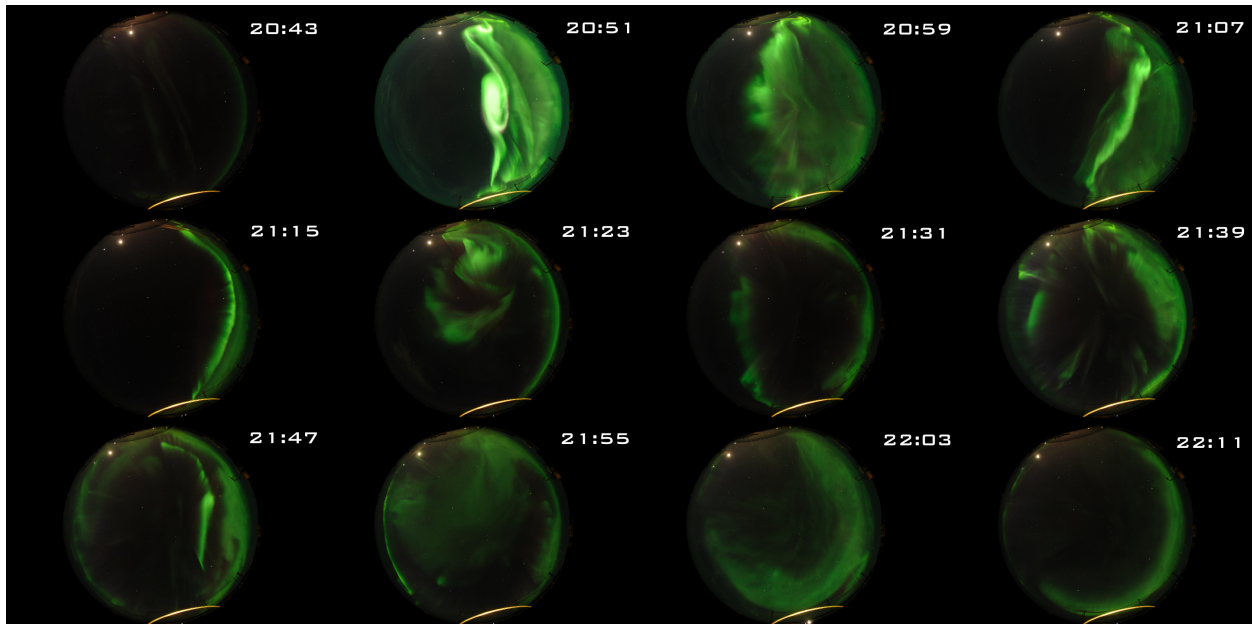


Figure 5: ALIS images compilation between 20:30 and 22:30 of March 28, 2012. The time in the upper right corner of each single picture represent the time it was taken.

References

- [1] Stenberg, G. (2012). *Practical 1. Aurora Borealis*. Swedish Institute of Space Physics, Kiruna, Sweden.
- [2] Kivelson M. G. and Russell C. T. (1996). *Introduction to Space Physics*. Cambridge University Press, Melbourne, Australia.