

# Level 2 Photometric Calibration for the LSST Survey

R. Lynne Jones<sup>1</sup>, Tim Axelrod<sup>2</sup>, Željko Ivezić<sup>1</sup>, David Burke<sup>3</sup>,  
James G. Bartlett<sup>4</sup>, Gurvan Bazin<sup>4</sup>, Guillaume Blanc<sup>4</sup>, Alexandre Boucaud<sup>4</sup>, Michel Crézé<sup>4</sup>,  
Cécile Roucelle<sup>4</sup>, Abhijit Saha<sup>5</sup>, J. Allyn Smith<sup>7</sup> Michael A. Strauss<sup>6</sup>, Peter Yoachim<sup>1</sup>  
on behalf of The Photometric Calibration Team

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## ABSTRACT

This document describes the photometric calibration procedure for LSST Data Release catalogs. This procedure will use specialized hardware, an auxiliary telescope, an atmospheric water vapor measurement system, and narrow-band dome screen illuminator, to measure the wavelength dependence of the atmospheric and hardware response functions, together with a self-calibration procedure that leverages multiple observations of the same sources over many epochs, to deliver 1%-level photometry across the observed sky.

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<sup>1</sup>University of Washington

<sup>2</sup>University of Arizona

<sup>3</sup>SLAC National Accelerator Laboratory

<sup>4</sup>APC, Université Paris Diderot

<sup>5</sup>NOAO

<sup>6</sup>Princeton University

<sup>7</sup>Austin Peay State University

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## 1. Introduction

LSST is required to deliver 1%-level photometry across the observed sky and under a wide range of observing conditions (0.5%-level for repeat observations of the same source). This represents about a factor of two improvement over the current state-of-the-art wide-field optical photometry delivered by SDSS under photometric conditions. This factor of two improvement will have a major impact on science deliverables because it implies that the error volume in the five-dimensional LSST color space will be over thirty times smaller than for SDSS-like photometry. This smaller error volume will improve source classification and the precision of quantities such as photometric redshifts for galaxies and photometric metallicity for stars. For example, a given spectral energy distribution (SED) corresponding to some galaxy type produces a line in the *ugrizy* multi-dimensional color space when redshifted, where the position of the galaxy along that line in *ugrizy* space is a function of redshift. Different galaxy SEDs produce lines that are often close to each other in *ugrizy* space and sometimes even cross. The smaller the error volume around an observed galaxy’s measured *ugrizy* colors, the smaller the number of different lines (thus, different SEDs) and different positions along the line (thus, different redshifts) which will be consistent with the measurement. The same conclusion is valid in the case of algorithms that estimate stellar effective temperature and metallicity, as well as any other model-based interpretation of measurements. Furthermore, the smaller error volume per source is advantageous even in the absence of any models. Two sources whose color differences produce a value of  $\chi^2$  per degree of freedom of 1, will have a  $\chi^2$  per degree of freedom of 4 when the errors are halved. In case of five degrees of freedom,  $\chi^2$  pdf  $> 4$  will happen by chance in only 0.1% of all cases. Therefore, the ability to reliably detect color differences between sources is a strong function of photometric errors.

The factor of two reduction in photometric error results from two major differences between LSST and SDSS. First, each source will receive hundreds of observations over the ten years of the LSST survey. These series of repeat observations will be used to self-calibrate

the photometric system across the sky and for each observation (akin, but not identical to, the uber-calibration procedure used by SDSS (Padmanabhan et al. 2008)), allowing LSST to operate in a wide variety of conditions. Secondly, the wavelength dependence of the hardware and atmospheric transmission response functions will be measured with auxiliary instrumentation on sufficiently fine angular and temporal scales to enable their explicit inclusion in the calibration procedure, rather than resorting to traditional approximations such as linear color terms. SNLS re-processing of CFHT Legacy Survey data found these color-dependent terms to be a significant contributor to the photometric calibration process (Regnault et al. 2009), on the level of a few to several percent. **TODO here on why it's hard - drift scan and all weather**

This document describes the calibration requirements and processes for LSST Data Release photometry. At each Data Release, there will be a complete recalibration of all data acquired to that point, on approximately an annual schedule. These data products are referred to as Level 2 Data Products. There will also be a real-time data calibration process, based on the best available set of prior calibrated observations, to provide real-time, but lower quality, photometry for quality assurance, generation of alerts, and other quantities appropriate for nightly data generation (aka Level 1 Data Products). The Level 1 photometric calibration is not discussed here.

Section 2 reviews the survey requirements for photometric calibration, while Section 3 describes the foundation of LSST’s calibration procedure, first motivating this procedure by describing the transmission of flux through the atmosphere and LSST system and then from the calibration point of view, trying to recreate the flux from the ADUs measured by the detector. Sections 4 and 5 describe those aspects in some detail. Section 6 describes how the LSST’s internal photometric scale is tied to external references. Section 7 describes the hardware required to realize the calibration process. Section 8 presents the error budget for each step of the calibration procedure. Section 9 describes how we will verify that the calibration system functions as designed, and meets the science requirements, first during the construction phase, and later during survey operations. Section 10 describes the implementation of the calibration process in software that will be part of LSST Data Management. Finally, Section 11 discusses the risks that remain in the implementation of the calibration process, and the steps we are taking to mitigate them.

## 2. Photometric Requirements

The LSST Science Requirements Document (SRD) provides a set of requirements on the annual Data Release (Level 2) photometry. These requirements are extended in the

XXX (OSS) to cover aspects which are too detailed for the SRD. based on measurements of bright, unresolved, isolated, non-variable objects from individual LSST visits. **TODO - address what these requirements actually apply to. MS stars vs other SEDs and reference App C. Also be explicit that calibration is based on objects with well-known SEDs - i.e. MS stars.** Bright implies that the measurement of the star's brightness is not dominated by photon statistics, approximately 1-4 magnitudes fainter than the saturation limit in a given filter. Isolated implies that the star can be successfully de-blended from background galaxies and other stars. Non-variable objects are intrinsically not variable; these will be identified in an iterative fashion from the many epochs of LSST observations. The SRD specifications are:

1. **Repeatability:** the median value of the photometric scatter for each star (the rms of calibrated magnitude measurements around the mean calibrated magnitude) shall not exceed 5 millimags in *gri*, 7.5 millimags in *uzy* for bright, unresolved, isolated, non-variable objects. No more than 10% of these objects should have a photometric scatter larger than 15 mmag in *gri*, 22.5 mmag in *uzy*. This specifies the distribution of photometric errors ( $\sigma$ ) and constrains both the repeatability of extracting counts from images and the ability to monitor (or model) the changes in the system response. It could be thought of as requiring the photometry of a single source to be consistent over time.
2. **Uniformity:** the rms of the internal photometric zeropoint error (for each visit) shall not exceed 10 millimags in *grizy*, 20 millimags in *uzy*, where the zeropoint for each visit is determined using bright, unresolved, isolated, non-variable sources. No more than 10% of these sources should be more than 15 mmag in *gri* or 22.5 mmag in *uzy* from the mean internal zeropoint. This places a constraint on the stability of the photometric system across the sky as well as an upper limit on various systematic errors, such as any correlation of photometric calibration with varying stellar populations (or colors). This makes the photometry of many sources comparable over the entire sky, and when combined with the previous requirement, creates a stable photometric system across the sky and over time, in a single filter.
3. **Band-to-band photometric calibration:** The absolute band-to-band zeropoint calibration for main sequence stars must be known with an rms accuracy of 5 millimags for any color not involving *u* band, 10 millimags for colors constructed with *u* band photometry. This places an upper limit on the systematic error in the measurement of the system throughput as a function of wavelength. This requirement ties photometric measurements in different filters together, enabling precise measurement of colors.

**4. Absolute photometric calibration:** The LSST photometric system must transform to an external physical scale (*e.g.* AB mags) with an rms accuracy of 10 millimags. This requirement not only ties LSST internal photometry to photometry obtained from other telescopes, using other photometric systems, but also ties LSST internal photometry to a real physical scale. This places a constraint on the upper limit of the systematic error in the measurement of the total system throughput. This final step enables LSST photometry to be compared directly to data from other telescopes or to models (*e.g.* such as determining the albedo of an asteroid with a known diameter).

Requirements 1 and 2 must be met by measuring and then correcting for changes in hardware and atmospheric transmission as a function of time, location in the sky or focal plane, and result in a relative calibration within a single filter. Requirements 3 and 4 require comparison of LSST measurements to externally calibrated spectrophotometric standards, providing a relative calibration from filter to filter as well as an absolute physical scale for the overall system. Performance of the LSST system regarding requirement 1 can be verified by simply measuring the rms of the calibrated magnitude measurements. Verification of requirement 2 is more complicated; in a simulated system it is simple to compare the (simulated, thus known) true magnitudes of the stars to the best-fit magnitudes produced after calibration. **TODO forward reference new section on verification - Tim to write** In operations, this will be verified using a combination of simulations, comparisons to known standards, and evaluation of science outputs such as stellar locus diagrams. These last two tests are also relevant to verifying the final two requirements, 3 and 4.

**TODO - Lynne (or someone) to verify consistency with new SRD reqs, \*after\* updating rest of doc and seeing effect of including new SEDs in SRD requirements. Plus need update of SRD to verify.**

### 3. The Photometric Calibration Process

In traditional photometric calibration, a set of standard stars are observed at a range of airmasses to calculate zeropoint offsets and (typically) a single color-dependent extinction curve per night. With care, this approach can deliver 1% photometry in stable photometric conditions. Such programs typically follow only a few objects, and devote roughly equal time to standards and program objects. This approach fails for a survey like LSST, for at least two reasons. First, from a calibration point of view, the very wide field and multiple detector array mean that effectively a large number of instruments must be calibrated rather than just one. Second, historical weather data from Cerro Pachón tells us only 53% of the available observing time can be considered photometric even at the 1–2% level. To take advantage

of the full 85% of the available observing time which is usable (total cloud extinction less than 1.5 magnitudes), and to reach the SRD specified requirements – 0.5% level photometric repeatability and 1% photometric uniformity – requires a new approach.

This new approach, variants of which are already in use at PanSTARRS and DES, *directly* measures the system throughput as a function of wavelength, focal plane position, and time. Further, the *normalization* of the throughput in each observation (the gray-scale zeropoint) and the *shape* of the throughput curve (the color dependent terms), are explicitly separated and measured with separate procedures for both the telescope system response and the atmospheric transmission. This calibration system requires various pieces of hardware to conduct these optimized measurements. We briefly describe them here, with full descriptions in Section 7:

- A dome screen projector designed to provide uniform ( $\sim 10\%$  variation) illumination across the field of view, while minimizing stray light. This projector system will have the capability to not only illuminate the screen with broadband white light, but also narrow-band light to measure the system response at individual wavelengths. The narrow-band light will be generated by a tunable laser, capable of producing light from 300 – 1100 nm and tunable in 1 nm increments. The brightness of the screen is measured with a NIST-calibrated photodiode, so that the relative intensity at different wavelengths is precisely determined.
- A 1.2-m auxiliary telescope with an  $R \approx 400$  spectrograph, located adjacent to the LSST itself. This auxiliary telescope will obtain spectra of a chosen set of atmospheric probe stars across the sky to determine an atmospheric absorption model.
- Water vapor monitoring system, consisting of a GPS system and a microwave radiometer attached copointed with the LSST telescope, and monitoring the same field of view. This supplements the auxiliary telescope spectra, which are unable to track the sometimes rapid variations of water vapor in time and space.

An overview of the entire calibration process, from science observation to calibrated photometric measurements, together with the required calibration data products is shown in Figures 4, 5, 6, and 7. Note that four classes of objects participate in the calibration process in different ways:

- Standard stars. These are stars whose absolute flux as a function of wavelength above the atmosphere is precisely known. This class contains only a few members, perhaps as few as ten. Their role is to enable the self calibration process to set absolute zeropoints for each band, and to allow testing of the SRD uniformity requirements. See Section 6.

- Calibration stars. These are stars that densely cover the sky, with typical spacings between stars of order 1 arcminute. Unlike standard stars, neither their SEDs nor their absolute fluxes are precisely known, and their standard magnitudes are determined by the self calibration process. They have been selected to be on the stellar main sequence, to be nonvariable, and to be relatively isolated so that their photometry is not degraded by crowding effects.
- Atmospheric probe stars. These are bright stars of known type, distributed roughly uniformly over the LSST sky, which yield high SNR spectra from the auxiliary telescope.
- Science objects. Calibration of science objects utilizes the results of processing the standards and the calibration stars, and the measurement of the system bandpass. If an SED is supplied for an object, an accurate standard magnitude can then be calculated.

The following section will provide a more in-depth overview of the calibration process. We will start with a review of what is physically happening to photons in their path toward the focal plane, and then outline how LSST will translate the measured ADU counts back to fluxes above the atmosphere.

We find it helpful to define four different magnitudes, and their associated fluxes:

- $m_b^{inst}$ , the instrumental magnitude.  $m_b^{inst} = -2.5 \log_{10}(C_b^{obs})$ , where  $C_b^{obs}$  are the instrumental counts (ADU) that are attributed to the object
- $m_b^{nat}$ , the natural magnitude. This is the magnitude in the AB system that would be measured for the object if it were measured through the actual normalized system bandpass,  $\phi_b^{obs}(\lambda)$ , at the top of the atmosphere. This bandpass varies from exposure to exposure. See equations 7 and 9.  $m_b^{nat} = m_b^{inst} + Z_b^{obs}$
- $m_b^{std}$ , the standard magnitude. This is the magnitude in the AB system that would be measured for the object if it were measured through the standard normalized system bandpass,  $\phi_b^{std}(\lambda)$ , at the top of the atmosphere. This bandpass is selected as part of the survey design, and does not vary. See equation 8.
- $m_b^{corr}$ , the SED corrected instrumental magnitude. This is the standard magnitude, but with an unknown gray zeropoint correction, which will be removed by self calibration. These magnitudes are the input to self calibration.  $m_b^{corr} = m_b^{inst} + \Delta m_b^{obs}$

These quantities are related through equations 15 and 15.

#### 4. From Flux to Counts

We first consider how the photons from an astronomical object make their way to the detector and are converted into counts (ADUs), paying attention to the various temporal or spatial scales for variability might arise in the LSST system to affect the final ADU counts.

Given  $F_\nu(\lambda, t)$  – the specific flux<sup>1</sup> (flux per unit frequency) of an astronomical object at the top of the atmosphere – at a position described by  $(alt, az)$ , the total flux from the object transmitted through the atmosphere to the telescope pupil is

$$F_\nu^{pupil}(\lambda, alt, az, t) = F_\nu(\lambda, t) S^{atm}(\lambda, alt, az, t), \quad (1)$$

where  $S^{atm}(\lambda, alt, az)$  is the (dimensionless) probability that a photon of wavelength  $\lambda$  makes it through the atmosphere,

$$S^{atm}(\lambda, alt, az, t) = e^{-\tau^{atm}(\lambda, alt, az, t)}. \quad (2)$$

Here  $\tau^{atm}(\lambda, alt, az)$  is the optical depth of the atmospheric layer at wavelength  $\lambda$  towards the position  $(alt, az)$ . Observational data (Stubbs et al. 2007; Burke et al. 2010) show that the various atmospheric components which contribute to absorption (water vapor, aerosol scattering, Rayleigh scattering and molecular absorption) can lead to variations in  $S^{atm}(\lambda, t)$  on the order of 10% per hour. Clouds represent an additional gray (non-wavelength dependent) contribution to  $\tau^{atm}$  that can vary even more rapidly, on the order of 2–10% of the total extinction at 1° scales within minutes (Ivezić et al. 2007).

Given the above  $F_\nu^{pupil}(\lambda, alt, az, t)$ , the total ADU counts transmitted from the object to a footprint within the field of view at  $(x, y)$  can be written as

$$C_b(alt, az, x, y, t) = C \int_0^\infty F_\nu^{pupil}(\lambda, alt, az, t) S_b^{sys}(\lambda, x, y, t) \lambda^{-1} d\lambda. \quad (3)$$

Here,  $S_b^{sys}(\lambda, x, y, t)$  is the (dimensionless) probability that a photon will pass through the telescope’s optical path to be converted into an ADU count, and includes the mirror reflectivity, lens transmission, filter transmision, and detector sensitivity. The term  $\lambda^{-1}$  comes from

<sup>1</sup>Hereafter, the units for specific flux (flux per unit are Jansky ( $1 \text{ Jy} = 10^{-23} \text{ erg cm}^{-2} \text{ s}^{-1} \text{ Hz}^{-1}$ )). The choice of  $F_\nu$  vs.  $F_\lambda$  makes the flux conversion to the AB magnitude scale more transparent, and the choice of  $\lambda$  as the running variable is more convenient than the choice of  $\nu$ . Note also, while  $F_\nu(\lambda, t)$  (and other quantities that are functions of time) could vary more quickly than the standard LSST exposure time of 15s, it is assumed that all such quantities are averaged over that short exposure time, so that  $t$  refers to quantities that can vary from exposure to exposure.

the conversion of energy per unit frequency into the number of photons per unit wavelength and  $b$  refers to a particular filter,  $ugrizy$ . The dimensional conversion constant  $C$  is

$$C = \frac{\pi D^2 \Delta t}{4gh} \quad (4)$$

where  $D$  is the effective primary mirror diameter,  $\Delta t$  is the exposure time,  $g$  is the gain of the readout electronics (number of photoelectrons per ADU count, a number greater than one), and  $h$  is the Planck constant. The wavelength-dependent variations in  $S_b^{sys}$  generally change quite slowly in time; over periods of months, the mirror reflectance and filter transmission will degrade as their coatings age. A more rapidly time-varying wavelength-dependent change in detector sensitivity (particularly at very red wavelengths in the  $y$  band) results from temperature changes in the detector, but only on scales equivalent to a CCD or larger. There will also be wavelength-dependent spatial variations in  $S_b^{sys}$  due to irregularities in the filter material; these are required by the camera specifications to vary (at a maximum) slowly from the center of the field of view to the outer edges, equivalent to a bandpass shift on the order of 1-2% of the effective wavelength of the filter. Wavelength-independent (gray-scale) variations in  $S_b^{sys}$  can occur more rapidly, on timescales of a day for variations caused by dust particles on the filter or dewar window, and on spatial scales ranging from the amplifier level, arising from gain changes between amplifiers, down to the pixel level, in the case of pixel-to-pixel detector sensitivity variations.

From equation 3 and the paragraphs above, we can see that the generation of counts  $C_b(alt, az, x, y, t)$  from photons is imprinted with many different effects, each with different variability scales over time, space, and wavelength. In particular the wavelength-dependent variability (bandpass shape) is typically much slower in time and space than the gray-scale variations (bandpass normalization). These different scales of variability motivate us to separate the measurement of the normalization of  $S_b^{sys}$  and  $S^{atm}$  from the measurement of the wavelength-dependent shape of the bandpass.

#### 4.1. Normalized bandpass response, $\phi_b(\lambda)$

This then leads us to introduce a ‘normalized bandpass response function’,  $\phi_b^{obs}(\lambda, t)$ , that represents the true bandpass response shape for each observation,

$$\phi_b^{obs}(\lambda, t) = \frac{S^{atm}(\lambda, alt, az, t) S_b^{sys}(\lambda, x, y, t) \lambda^{-1}}{\int_0^\infty S^{atm}(\lambda, alt, az, t) S_b^{sys}(\lambda, x, y, t) \lambda^{-1} d\lambda}. \quad (5)$$

Note that  $\phi_b$  only represents *shape* information about the bandpass, as by definition

$$\int_0^\infty \phi_b(\lambda) d\lambda = 1. \quad (6)$$

Using  $\phi_b^{obs}(\lambda, t)$  we can represent the (true, total) in-band flux of an object for each observation as

$$F_b^{obs}(t) = \int_0^\infty F_\nu(\lambda, t) \phi_b^{obs}(\lambda, t) d\lambda, \quad (7)$$

where the normalization of  $F_b(t)$  corresponds to the top of the atmosphere. Unless  $F_\nu(\lambda, t)$  is a flat ( $F_\nu(\lambda) = \text{constant}$ ) SED,  $F_b^{obs}$  will vary with changes in  $\phi_b^{obs}(\lambda, t)$  due simply to changes in the bandpass shape, such as changes with position in the focal plane or differing atmospheric absorption characteristics, *even if the source is non-variable*.

To provide a reported  $F_b^{std}(t)$  which is constant for non-variable sources, we also introduce the ‘standardized bandpass response function’,  $\phi_b^{std}(\lambda)$ , a curve that will be defined before the start of LSST operations (most likely during commissioning).  $\phi_b^{std}(\lambda)$  represents a typical hardware and atmospheric transmission curve, minimizing the difference between  $\phi_b^{obs}(\lambda, t)$  and the standardized reported bandpass. Now,

$$F_b^{std}(t) = \int_0^\infty F_\nu(\lambda, t) \phi_b^{std}(\lambda) d\lambda, \quad (8)$$

is a constant value for non-variable sources.

Magnitudes provide an easy way to conceptualize the relationship between  $F_b^{obs}$  and  $F_b^{std}$ , provided that we define a ‘natural magnitude’

$$m_b^{nat} = -2.5 \log_{10} \left( \frac{F_b^{obs}}{F_{AB}} \right) \quad (9)$$

where  $F_{AB} = 3631$  Jy. The natural magnitude, like  $F_b^{obs}$  will vary from observation to observation as  $\phi_b^{obs}(\lambda, t)$  changes, even if the source itself is non-variable. The natural magnitude can be translated to a ‘standard magnitude’,  $m_b^{std}$ , as follows:

$$m_b^{nat} = -2.5 \log_{10} \left( \frac{F_b^{obs}}{F_{AB}} \right) \quad (10)$$

$$= -2.5 \log_{10} \left( \frac{\int_0^\infty F_\nu(\lambda, t) \phi_b^{obs}(\lambda, t) d\lambda}{F_{AB}} \right) \quad (11)$$

$$= -2.5 \log_{10} \left( \left( \frac{\int_0^\infty F_\nu(\lambda, t) \phi_b^{obs}(\lambda, t) d\lambda}{\int_0^\infty F_\nu(\lambda, t) \phi_b^{std}(\lambda, t) d\lambda} \right) \left( \frac{\int_0^\infty F_\nu(\lambda, t) \phi_b^{std}(\lambda, t) d\lambda}{F_{AB}} \right) \right) \quad (12)$$

$$m_b^{nat} = \Delta m_b^{obs} + m_b^{std} \quad (13)$$

$$\Delta m_b^{obs} = -2.5 \log_{10} \left( \frac{\int_0^\infty F_\nu(\lambda, t) \phi_b^{obs}(\lambda, t) d\lambda}{\int_0^\infty F_\nu(\lambda, t) \phi_b^{std}(\lambda, t) d\lambda} \right) \quad (14)$$

where  $\Delta m_b^{obs}$  varies with the *shape* of the source spectrum,  $f_\nu(\lambda, t)$  and the *shape* of the bandpass  $\phi_b^{obs}(\lambda, t)$  in each observation. Note that  $\Delta m_b^{obs} = 0$  for flat (constant) SEDs, as

the integral of  $\phi_b(\lambda)$  is always one. For non-variable sources,  $m_b^{std}$  will be non-variable as it represents the throughput in a standardized bandpass,  $\phi_b^{std}(\lambda)$ .

The natural and standard magnitudes can be tied back to the counts produced by the system by adding the correct zeropoint offsets. As  $\Delta m_b^{obs}$  removes all wavelength dependent variations in  $m_b^{std}$ ,

$$m_b^{std} = m_b^{inst} - \Delta m_b^{obs} + Z_b^{obs} \quad (15)$$

$$\equiv m_b^{corr} + Z_b^{obs} \quad (16)$$

where

$$m_b^{inst} = -2.5 \log_{10}(C_b^{obs}) \quad (17)$$

The zeropoint correction here,  $Z_b^{obs}$ , contains only gray-scale *normalization* effects, such as variations due to the flat field or cloud extinction. The SED corrected magnitude,  $m_b^{corr}$ , is the input to the Self Calibration block in Figure 4.

#### 4.2. Perturbations to the System Bandpass

The normalized system bandpass will vary from exposure to exposure due to a myriad of effects. The major ones are:

- Atmospheric transmission variations. The major wavelength-dependent sources of atmospheric absorption, Rayleigh scattering from molecules, molecular oxygen lines, ozone, water vapor, and aerosols, all vary in time and space, with widely varying temporal and spatial scales. Even in the absence of intrinsic variation in the atmosphere, of course, the bandpass varies due to changing airmass. This latter effect is compensated for in traditional photometric calibration, but the others are generally ignored
- Long term variations in the throughput of the optical path due to contamination
- Changing detector quantum efficiency, particularly in the y-band, due to varying focal plane temperature
- Shifts in filter position with respect to the system optical axis, due to positioning jitter and gravity sag

It bears repeating that every perturbation in general affects both the zeropoint, through the gray component of the perturbation, and the shape of the system bandpass, through the

wavelength-dependent component. The gray component is removed by the self calibration process, while the wavelength-dependent component must be separately characterized and removed (see Figure 4). We are concerned here only about the latter effect, and discuss it for each of the above categories.

### 4.3. Effects of Airmass Variation

The effects of airmass variation on photometry is well known to all photometrists. In fact, it is the only effect which is always accounted for in photometric calibration, and quite often the only effect. Figure ?? shows the effects of variation over the full airmass range expected for the LSST survey. There is a more subtle effect, however, which is important because of the LSST’s large field of view: the airmass can vary significantly from one side of the field to the other. For example, if the field center is airmass 2.1, the airmass varies from 1.98 to 2.22 across the field. Figure 8 shows the error that would be made in ignoring this effect. This requires us to maintain an atmospheric model which can be interpolated to any position in the focal plane.

### 4.4. Effects of Atmospheric Variations

The main components of wavelength-dependent atmospheric extinction are Rayleigh scattering from molecules; oxygen molecular lines; water vapor; ozone; and aerosols. Rayleigh scattering and molecular oxygen extinctions are directly proportional to barometric pressure at the surface. As is well known from looking at surface pressure charts, away from weather fronts the pressure varies significantly on timescales of hours or more and spatial scales of hundreds of km. Variations can be much more rapid in the vicinity of fronts. In any case, this component can be compensated for very accurately just by measuring the barometric pressure, and we do not further consider it here. The other components are not so easily measured, and in the case of water vapor and aerosols, can display complex patterns of variability.

Figures 9, 10, and 11 give three different perspectives on historical water vapor variability in the vicinity of CP. There are broad patterns in space (the E-W gradient), and in time (the regular seasonal variations). There are large variabilities on top of these, amounting to several mm of PWV, which can occur in hours or less. The biggest effect of water vapor is in the y-band, since it contains a strong water band (Figure ??). Figure 12 shows the change in natural magnitude when the PWV is varied from 1mm to 6mm for a set of Kurucz

stars. The 4mm variation in PWV shown in Figure 10 would lead to a roughly +/-3mmag color-dependent scatter in calibration of y photometry, if not compensated for.

While we do not have aerosol variability data for Cerro Pachon (CP), data from CASLEO, a site at 2550m in Argentina ((?)) serves as a reasonable proxy. Figure 13 shows the time history of aerosol optical depth at 675nm over a period of roughly a year. The spiky nature of the data is notable, suggesting rapid variations on timescales of perhaps a few hours. Figure 14, which shows a time-altitude profile of aerosol variations at a site in the US, backs up this impression. Although the site is at much lower altitude than CP, there are nonetheless significant variation of aerosol extinction at altitudes above 3000m on time scales well under an hour, and we should expect similar variations at CP. Determining the effects of aerosol variations on the system bandpass is more complex than for other atmospheric components, because the spectral shape of the aerosol extinction varies as well as its magnitude. Figure 15 shows the change in natural magnitude when the aerosol optical depth is varied from 0.04 mag to 0.16 mag, roughly the range of the CASLEO data, while keeping a constant spectral index of  $\alpha = -1.7$ . Unlike water vapor, which affected the reddest bands most strongly, aerosols affect the bluest bands the most. The effects in the u-band range from -5 mmag to +35 mmag, strongly dependent on star color.

Ozone is dominantly a stratospheric component of the atmosphere, and is routinely monitored by satellite. There is little evidence for variations on short time or space scales, with most variation on a seasonal scale. Figure 16 shows the overall variability of ozone at CP over a roughly 8 year period. Figure 17 shows the changes in the natural magnitudes for a variation in ozone by 50 Dobson units. The effect is very small, except in the u-band, where it has a noticeable effect on red stars.

To summarize the effects of wavelength-dependent atmospheric variations, they are sufficiently large, and occur sufficiently rapidly, that they must be corrected by an atmospheric model with fidelity substantially greater than the traditional photometrist's extinction model. The combination of the auxiliary telescope and the water vapor monitoring system will supply the data required for construction of these atmospheric models. It is also worth emphasizing that correcting for these effects requires not only an accurate bandpass model, but also SEDs for the objects being calibrated. Calibration stars will be picked from well defined stellar populations to minimize SED uncertainty. SED determination for arbitrary science objects, such as supernovae and galaxies, will be more challenging, and may limit the accuracy of their photometric calibration.

#### 4.5. Throughput Variations Due to Contamination

#### 4.6. Variations in Detector Quantum Efficiency

#### 4.7. Throughput Variations Due to Filter Position Shifts

#### 4.8. Putting it All Together

Examples of the  $\Delta m_b^{obs}$  due to variations in the shape of the hardware and atmospheric response curves are shown in Figure 18 and Table 2. Two main sequence stellar models (Kurucz 1993) – one with temperature 35000K (blue) and one 6000K (red) – were combined with three different atmospheric response curves (with airmass  $X=1.0$  with minimal H<sub>2</sub>O vapor,  $X=1.2$  with a nominal amount of H<sub>2</sub>O (the ‘standard’), and  $X=1.8$  with a large H<sub>2</sub>O vapor content) and two different hardware response curves (one ‘standard’ and one shifted in wavelength by 1%) to illustrate the resulting changes in observed natural magnitudes. In Figure 19, the  $X = 1.8$  atmospheric response is combined with a 1% shift (the maximum allowed in the filter manufacturing specification from center to edge) in filter bandpass, thus altering the hardware response, for many main sequence Kurucz models spanning a range of  $g - i$  colors; the resulting changes in natural magnitudes are plotted. These examples demonstrate that the scatter in natural magnitudes induced by expected atmospheric and hardware transmission curve shape changes alone (without any gray-scale changes) can be much larger than the SRD repeatability requirements would permit. Roughly speaking, these effects reach a level of about 50 mmag. This suggests that our measurement-based model of wavelength-dependent effects must be accurate at the 90-95% level. This issue is discussed in detail in Section 8

### 5. From Counts to Flux

The previous section laid out the origins of ADU count variability from one observation to another. Now we will consider how we can, in practice, acquire the information necessary to convert a particular observed ADU count to a measurement of  $F_\nu(\lambda, t)$  above the atmosphere for a particular object. This requires measuring and then compensating for the variations in  $S^{atm}(\lambda, alt, az, t)$  and  $S_b^{sys}(\lambda, x, y, t)$ . Let us first consider measurement of the variations in the hardware throughput curve,  $S_b^{sys}(\lambda, x, y, t)$ .

Table 1::  $\Delta m_b^{obs}$  due to variations in system and atmospheric bandpass shape (see also Fig 18). The first two rows show the baseline ('standard') magnitude of the star. All other rows show the *change* in magnitude (in mmag) due to the variations listed at left. Any value larger than 5 mmag would be larger than the RMS scatter allowed by the SRD. *TODO color-code values larger than 5 mmag*

| Bandpass                   | star | <i>u</i> (mag)    | <i>g</i>   | <i>r</i>   | <i>i</i>   | <i>z</i>   | <i>y</i>   |
|----------------------------|------|-------------------|------------|------------|------------|------------|------------|
| Std (X=1.2) atm, std sys   | red  | 21.472            | 20.378     | 20.000     | 19.911     | 19.913     | 19.913     |
| Std (X=1.2) atm, std sys   | blue | 19.102            | 19.503     | 20.000     | 20.378     | 20.672     | 20.886     |
|                            |      | $\Delta u$ (mmag) | $\Delta g$ | $\Delta r$ | $\Delta i$ | $\Delta z$ | $\Delta y$ |
| Std (X=1.2), +1% sys shift | red  | -31               | -22        | -8         | -2         | 1          | 1          |
| Std (X=1.2), +1% sys shift | blue | 9                 | 17         | 20         | 20         | 16         | 16         |
| X=1.0, std sys             | red  | 7                 | 2          | 0          | 0          | -0         | -1         |
| X=1.0, std sys             | blue | -3                | -1         | -1         | -0         | 1          | -4         |
| X=1.0, +1% sys shift       | red  | -24               | -20        | -8         | -1         | 1          | 0          |
| X=1.0, +1% sys shift       | blue | 7                 | 16         | 19         | 20         | 18         | 12         |
| X=1.8, std sys             | red  | -21               | -10        | -2         | -0         | 0          | 1          |
| X=1.8, std sys             | blue | 8                 | 8          | 4          | 2          | -1         | 6          |
| X=1.8, +1% sys shift       | red  | -50               | -30        | -10        | -2         | 1          | 2          |
| X=1.8, +1% sys shift       | blue | 16                | 24         | 24         | 22         | 15         | 22         |

### 5.1. Accounting For Finite PSF Width

The discussion in Section 4 implicitly assumed that the flux from an object is sampled at a single point in the focal plane. We must now extend the formalism to account for the fact that the system has a PSF with finite width, so that the flux from an object is spread over multiple pixels, each of which may differ slightly in response from its neighbors. Let us use  $i$  as an index over the pixels in the PSF for a particular object, and let  $w_i$  be the integral of the PSF over the area of pixel  $i$ , with  $\sum_i w_i = 1$ . To make the notation compact, define

$$q_{b,i}(\lambda) = S^{atm}(\lambda, alt, az, t) S_b^{sys}(\lambda, x_i, y_i, t) \lambda^{-1} \quad (18)$$

Further, define the SED of the object,  $f_\nu(\lambda)$  by

$$F_\nu(\lambda) = F_b f_\nu(\lambda) \quad (19)$$

where  $\lambda$  is within band b, and

$$\int_b f_\nu(\lambda) d\lambda = 1 \quad (20)$$

Then the total number of counts from the object,  $C_b^{inst}$  is

$$C_b^{inst} = F_b \sum_i w_i \int q_{b,i}(\lambda) f(\lambda) d\lambda \quad (21)$$

For compactness, let's define

$$Q_{b,i} = \int q_{b,i}(\lambda) f_\nu(\lambda) d\lambda \quad (22)$$

Since we are following the path from counts to flux, we are interested in

$$F_b = \frac{C_b^{inst}}{\sum_i w_i Q_{b,i}} \quad (23)$$

At this point, we notice that  $Q_{b,i}$  is close to what we mean by a 'flat'. Let us make it identical, by factoring out the normalization, which is not directly measured:

$$Q_{b,i} = N Q_{b,i}^{norm} \quad (24)$$

where  $N$  is chosen to make  $\max_i(Q_{b,i}^{norm}) = 1$ .  $N$  will become the zeropoint of the exposure:  $Z_b^{obs} = -2.5 \log_{10}(N)$ .

There is a problem with using this flat, however: At the time we process the pixels in Flat SED Calibration (Figure 4), we are unable to calculate  $Q_{b,i}$  because we are (probably) ignorant of both  $f_\nu(\lambda)$  and the atmospheric bandpass for the observation,  $S_b^{atm}(\lambda)$ . Even if we knew that information, this prescription would be unworkable because the flat would need to be constructed with an SED that varies rapidly over the focal plane. So we must accept our inability to work with  $Q_{b,i}$  when processing the pixels, and must content ourselves with a broadband flat measured from the dome flatfield system.

$$Q_{b,i}^{BB} = \int q_{b,i}^{sys}(\lambda) f_\nu^{BB}(\lambda) d\lambda \quad (25)$$

where  $f_\nu^{BB}(\lambda)$  is an SED we are free to choose in constructing the broadband flat. Note that  $q_{b,i}^{sys}(\lambda) = S_b^{sys}(\lambda, x_i, y_i, t) \lambda^{-1}$  includes only the system response, not the atmosphere.

We can now write

$$F_b = \frac{1}{N} \left[ \frac{C_b^{inst}}{\sum_i w_i Q_{b,i}^{BB}} \right] \left[ \frac{\sum_i w_i Q_{b,i}^{BB}}{\sum_i w_i Q_{b,i}} \right] \quad (26)$$

(normalize flat, bring in Zp) Expressed in magnitudes, this becomes

$$m_b^{nat} = m_b^{inst} + Z_p - 2.5 \log_{10} \left( \left[ \frac{\sum_i w_i Q_{b,i}^{BB}}{\sum_i w_i Q_{b,i}^{norm}} \right] \right) \quad (27)$$

This may not at first seem like an advance over equation 23, but in fact it is: The first term is just what we usually mean by measuring the flux in a flattened image. It is done at the pixel level in the usual way, without any dependence on  $f(\lambda)$ . We can at least imagine calculating the second, correction, term in exactly the way it is written. We do not need to actually access the image pixels, but do need to know the PSF and the object SED, and we need to construct at least a localized flat from the object SED. We reserve this as a future option in calibration processing, because it may reduce the systematics floor in the instrumental magnitudes. For the present, however, we make the assumption which is univerally made today in doing photometry, albeit almost always implicitly: The ratio of the true flat to the broadband flat is independent of the pixel within the PSF footprint, so that:

$$m_b^{nat} = m_b^{inst} + Z_p - 2.5 \log_{10} \left( \frac{Q_b^{BB}}{Q_b^{norm}} \right) \quad (28)$$

Note that the correction term must be applied in the "SED Correction" block of Figure 4. The required information to compute it is not available earlier. We now turn to the actual measurement of  $q_{b,i}$ , and the construction of  $Q_{b,i}^{BB}$

## 5.2. Measuring the Hardware Response

To measure the wavelength-dependent hardware response curve as a function of position in the focal plane, we will use a dome-screen system that is capable of producing narrow-band light over a range of wavelengths, producing a data cube of 'narrow-band flat fields'. A similar approach has already been employed at PS-1 ((Stubbs & Tonry 2012); (Tonry et al. 2012)), and at DES ((Marshall et al. 2013)) After applying the monochromatic illumination correction (see Section 5.2.1) and applying the normalization at each wavelength from the

photodiode monitors (see Section XX), this data cube records (up to an overall normalization constant)  $S_b^{sys}(\lambda, t)$  at each  $x, y$  location in the focal plane. This data cube of narrow-band flats is used to construct a synthetic broadband flat field ('BBF') by combining the individual narrow-band flats according to a chosen spectral energy distribution.

As discussed in the previous section, the choice of the SED used to construct the broadband flats is largely arbitrary. It defines a reference SED  $f_\nu^{BB}(\lambda)$ , which must later be corrected for as shown in equation 28 when both the true  $\phi_b^{obs}(\lambda)$  and the actual object SED,  $f_\nu(\lambda)$  are known.

Given the arbitrariness of  $f_\nu^{BB}(\lambda)$ , we must choose on grounds of convenience. At least two choices suggest themselves:

- Choose  $f_\nu^{bbf}(\lambda) = const.$  This has the virtue of maximal simplicity.
- Choose  $f_\nu^{bbf}(\lambda)$  to be in some sense an average SED for all objects in the survey, propagated through an average atmosphere, such that the average value of the BBF correction is minimized.

We choose the first of these, and subsequently refer to the BBF as the FSF (flat spectrum flat). This choice is bound to have unaccounted for impact on calibration accuracy, however, because the above formalism does not include the propagation of errors. We can readily change the prescription during construction or commissioning.

Generation of the entire data cube of narrow-band flats is too time-consuming to complete on a daily basis. Instead, the full narrow-band flat field scan will only be repeated on a time interval adequate for measurement of the more slowly variable components of  $S_b^{sys}(\lambda, t)$ , approximately monthly (but to be determined during commissioning). Since the system response can change on shorter timescales, principally due to a changing population of dust particles on optical surfaces, we correct the FSF nightly by multiplying by the ratio of two broad-band flats, one at the current epoch and one at the reference epoch when the narrow-band flats were acquired.

### 5.2.1. Determining the Illumination Correction

As mentioned above, before dome flats (either broadband or narrow-band) can be used to measure  $S_b^{sys}$ , they must be modified to correctly produce *photometrically* uniform measurements of a collimated source across the field of view. This correction is called the 'illumination correction'. The illumination correction must correct the observed flat fields

for effects resulting from non-uniform illumination of the dome screen, for ghosting caused by internal reflections in the camera, and for the presence of stray or scattered light arriving in the focal plane on paths other than the direct image path (such as light bouncing from the dome floor or glinting off a filter holder). Figure 20 gives a perspective on the difference between the direct, total, and ghost flatfield illumination from a realistic model of the LSST telescope, camera, and domescreen illumination system.

The illumination correction is difficult to measure directly. To do so would require a collimated light source with diameter small compared with the telescope pupil that is able to scan in pupil position, angle, and wavelength. Aside from the engineering challenges in producing such a light source in the dome, the resulting 5-dimensional data cube would require an inordinate amount of time to collect, and we have deemed this approach impractical. Instead, we plan to combine data from several paths to generate an illumination correction which is consistent with all of them:

- Detailed monte carlo optics models, such as FRED
- Laboratory measurements of the as-built camera with the CCOB (see Appendix X). This provides the detailed data cube envisioned above, but only for the camera in isolation. This is sufficient to characterize ghosts, but not scattered light in the full system.
- Raster scans of star fields in photometric conditions. The star fields will be chosen to be dense, without compromising photometry due to crowding, and to contain a wide range of colors.
- Second order corrections from running self calibration. Self calibration is particularly capable at determining corrections which are position dependent but wavelength independent, such as arise from nonuniform illumination of the dome screen (see Figure 22).

This process will consume significant amounts of time in the lab, time on the sky, and time in analysis. Fortunately, the illumination correction is expected to be stable with time and will be remeasured only when components in the optical path of the telescope are altered.

Because the illumination correction is wavelength dependent (highly so near the edges of the filter bandpasses), its effect on photometry is dependent on the SED of the object in question. Figure 21 shows this effect as a function of both radius in the focal plane and star color, for a set of Kurucz main sequence stars. The effects are clearly too large to ignore, or they would overwhelm the error budget. This is discussed further in Section 8. It seems

likely that the quality of the illumination corrections is currently limiting the photometric accuracy of large surveys, and we expect the LSST treatment of the illumination correction to yield significant improvement.

### 5.3. Measuring the Atmospheric Transmission

Next, considering  $S^{atm}(\lambda, alt, az, t)$ , we will again separate the measurement of the shape of the atmospheric response and the measurement of normalization of the transmission. The currently available data is still incomplete, but suggests that the wavelength-dependent variations in  $\phi^{atm}(\lambda, t)$  change smoothly over spatial scales larger than the field of view and over several minutes. By using an auxiliary telescope equipped with a spectrograph to observe bright stars with known SEDs, we can measure atmospheric absorption at a variety of positions in the sky every 5–10 minutes throughout the night. These observations are used as constraints for MODTRAN atmospheric models, generating representations of the atmospheric throughput in the form of a set of absorption components as a function of  $alt, az, t$ . These components can be interpolated in time and space to generate a wavelength-dependent atmospheric absorption profile,  $\phi_b^{atm}(\lambda, alt, az, t)$ , for each observation.

This approach has been put into practice by Burke et al ((Burke et al. 2010)). Probe stars were selected in the range  $9 < V < 12$ , and spectra were taken on a 1.5m telescope with a  $R \approx 400$  resolution spectrograph. Exposure times were 2 to 4 minutes. The observing pattern is shown in Figure 23. The atmospheric model was a simple linear combination of templates generated by Modtran4 for molecular scattering, molecular absorption, aerosols, ozone, and water vapor. The models, together with some factors to account for the spectrograph efficiency as a function of wavelength, were fit to all the observed spectra simultaneously. Figures x and y show some results. The fit to the individual spectra are impressively good, and the coefficients in the resulting atmospheric models show strong variation, as expected for the highly variable nights for which the data was obtained. Traditional extinction plots (Figure z), show behavior that would be well fit by traditional calibration in the r- and i-bands, together with behavior in the z- and y-bands that would not. In subsequent development of the approach, Burke et al ((?)) obtained probe star spectra while simultaneously imaging over a wide field - roughly comparable to LSST operations. Unlike LSST, however, the imaging exposures were managed to keep the pixel coordinates of a given star nearly constant for all exposures. The full calibration process on the imaging data, including both atmospheric fitting and gray extinction determination through self calibration, yielded repeatability of approximately 8 mmag - good enough to meet the SRD minimum requirement for repeatability.

Using MODTRAN we can generate atmospheric transmission profiles at a variety of airmasses for each of these major sources of atmospheric extinction – molecular (Rayleigh) scattering, aerosol (Mie) scattering, and molecular absorption from each of O<sub>3</sub>, H<sub>2</sub>O, and combined O<sub>2</sub>/trace species, as is shown in Figure 30 for a standard atmospheric composition (the 1976 US Standard). These profiles capture the wavelength dependence of each component individually, over a grid of airmasses, and can be used as templates to generate new atmospheric transmission curves for any desired atmospheric composition as follows:

$$\begin{aligned}
 S^{fit}(alt, az, t, \lambda) = & e^{-\tau_{aerosol}(alt, az, t, \lambda) X} \\
 & \times (1 - C_{mol}(BP(t)/BP_o) A_{Rayleigh}(X, \lambda)) \\
 & \times (1 - \sqrt{C_{mol}(BP(t)/BP_o)} A_{O_2}(X, \lambda)) \\
 & \times (1 - C_{O_3}(t) A_{O_3}(X, \lambda)) \\
 & \times (1 - C_{H_2O}(alt, az, t) A_{H_2O}(X, \lambda)). \tag{29}
 \end{aligned}$$

The  $A_{Rayleigh/O_2/O_3/H_2O}$  functions are absorption templates (i.e. 1 minus the transmission profiles from the MODTRAN models), the  $C_{mol,O_3,H_2O}$  are coefficients describing the composition of the atmosphere together with  $\tau_{aerosol}$ , and  $BP(t)$  is measured. An example of an atmosphere generated in this fashion is shown in Figure 31, demonstrating that this method can be used to generate an atmosphere at any airmass for any composition desired, without needing to generate a full MODTRAN model.

With this capability, we can fit the auxiliary telescope spectroscopic data taken throughout the night for the values of  $C_{mol,O_3,H_2O}$ , increasing our SNR for these coefficients by modeling their expected behavior over time and across the sky as detailed in ?? above. The Rayleigh scattering and molecular absorption due to O<sub>2</sub> and other trace species are fit with a single coefficient,  $C_{mol}$ , which simply scales the MODTRAN templates to the appropriate level for Cerro Pachon, and then only change with the barometric pressure ( $BP$ ). The O<sub>3</sub> absorption is fit with a single  $C_{O_3}$  value for each night, as it is not expected to vary more than 5-10% within a night. The aerosol absorption, as it is expected to have a small spatial variation across the sky, is modeled as

$$\tau_{aerosol}(alt, az, t, \lambda) = (\tau_0 + \tau_1 \text{EW} + \tau_2 \text{NS}) \left( \frac{\lambda}{\lambda_0} \right)^\alpha, \tag{30}$$

where EW and NS are defined as EW = cos(alt)sin(az), NS = cos(alt)cos(az), projections of the telescope pointing in the EW/NS directions. Single values of  $\tau_0$ ,  $\tau_1$ ,  $\tau_2$  and  $\alpha$  are fit for each night of observing, with  $\tau_1$  and  $\tau_2$  likely to be very small (Burke et al. 2010). The H<sub>2</sub>O absorption is likewise expected to show spatial variation, but also time variability, and is modeled as

$$C_{H_2O}(alt, az, t) = C_{H_2O}(t) + \frac{dC_{H_2O}}{d\text{EW}} \text{EW} + \frac{dC_{H_2O}}{d\text{NS}} \text{NS} \tag{31}$$

using a constant spatial EW and NS gradient per night and a  $C_{H_2O}(t)$  that is fit to each auxiliary telescope measurement (and interpolated between these times).

The coefficients  $C_{mol/O_3/H_2O}$  and  $\tau^{aerosol}$  will be determined using spectra of bright stars obtained from the 1.2-m LSST auxiliary telescope. The auxiliary telescope will be equipped with a modest resolution ( $R \sim 400$ ) spectrograph, sufficient to capture the signatures of the atmospheric extinction components, and covering the entire wavelength range of LSST ( $300 < \lambda < 1100$  nm) in each exposure. The stars observed with the auxiliary telescope must be bright ( $r < 12$ ) and ideally either white dwarfs or F stars – stars with relatively simple and well-understood SEDs to minimize confusion with the atmospheric extinction. By observing the same grid of stars on multiple nights, even if the SEDs are not well determined initially, they can be bootstrapped from the many epochs of data.

#### 5.4. Finding the Zero Points: Self Calibration

In order to correct for the more rapid gray-scale variations in the relative normalization of  $S^{atm}(alt, az, t)$  due principally to cloud extinction, we must use the observations of calibration stars in the images themselves, as observations ((?), (Ivezic et al. 2007)) suggest that cloud extinction can vary by 0.01 magnitudes on the scale of a CCD on timescales as fast as a few minutes. This ‘self-calibration’ procedure could be thought of as creating a massive calibration ‘standard’ star catalog, where the calibration stars are a selected set of the non-variable, main-sequence stars in the science images; their main difference from true standards is that the true magnitudes of the calibration stars have to be bootstrapped from the many different observations of the survey, and their SEDs need to be inferred from multicolor photometry. For every calibration star, corrections for  $\phi_b^{sys+atm}(\lambda, t)$  must be applied to produce a standardized magnitude,  $m_b^{std}$ , then in the self-calibration procedure we minimize the difference between the standardized magnitude and a model magnitude,

$$\chi^2 = \sum_{ij} \left( \frac{m_{b,ij}^{std} - m_{b,ij}^{model}}{\sigma_{b,ij}^{std}} \right)^2 \quad (32)$$

where the model magnitude is derived from the best-fit ‘true’ magnitude of the calibration star and a model describing how we expect the magnitude to vary from observation to observation. In the simplest self-calibration plan, this model simply consists of a normalization constant (zeropoint offset) for a ‘patch’ equivalent to the size of a CCD,

$$m_{b,ij}^{model} = m_{b,i}^{best} - \delta z_{b,j}. \quad (33)$$

This produces best-fit magnitudes for the calibration star catalog as well as zeropoint offsets (normalization constants) for each CCD in every observation, allowing us to correct for atmo-

spheric extinction on the scale of a CCD. By adopting a more complex model, this procedure can also correct for variations in the relative normalization of the total system throughput beyond those contributed by cloud extinction (such as remaining errors in the illumination correction for the broadband and narrow-band flat fields), but is generally limited by the number of stars and number of observations of each star that are obtained. A CCD size patch provides several hundred stars calibration stars per patch, allowing good signal to noise when determining cloud extinction which varies from observation to observation. This is similar in nature to the ubercal method applied to SDSS in Padmanabhan et al. (2008), and more recently DLS (Wittman et al. 2012) and PanSTARRS-1 (Schlafly et al. 2012).

**TODO:** Need to discuss healpix implementation, incorporation of standards, etc

Repeating Equation 15 above, adjusting  $obs$  indexes to  $meas$  to reflect the difference between the true and measured quantities,

$$m_b^{std} = m_b^{inst} - \Delta m_b^{meas} + Z_b^{meas} \quad (34)$$

we can relate the terms in this equation to the corrections just described above.  $\Delta m_b^{meas}$  originates from the difference between  $\phi_b^{meas}(\lambda, t, x, y)$  and  $\phi_b^{std}(\lambda)$  convolved with the source SED, and thus it depends on the shape of the total system response as well as the shape of the source SED.  $\Delta m_b^{meas}$  will be calculated by combining a series of model SEDs with  $\phi_b^{meas}(\lambda, t, x, y)$  at various locations in the focal plane, creating a lookup table of values to apply to measured magnitudes. The  $Z_b^{meas}$  zeropoint offset comes from any normalization constants generated by the self-calibration procedure (in the simple model, just the  $\delta z_{b,j}$  in equation 33 above).

These standard magnitudes are calibrated for variations in the observed bandpass shape (where applicable) and relative normalization, thus are directly comparable from one observation to the next. However, they are not yet tied to an external physical scale or from one filter band to another, and thus only define an internally calibrated LSST magnitude in a particular filter.

To fulfill SRD requirements 3 and 4, these internally calibrated natural magnitudes must also be tied from one filter band to another, and then tied to an absolute external physical scale. For this, a further set of measurements is needed. In all filters, a set of spectrophotometric standards must be observed, and calibrated using the steps described above. Then the known SED is combined with the standard bandpass shape to generate synthetic color photometry. The synthetic colors are then compared with the calibrated measured standard magnitudes to calculate  $\Delta_{b-r}$ , the corrections needed to tie measurements in each filter together (referenced to  $r$  band). At this point, only one final measurement is

necessary to tie the entire system to an external physical scale: an  $r$  band LSST natural magnitude measurement of an absolutely calibrated source on a photometric night. Although in theory these last two steps could be done with a single externally calibrated object, on a single photometric night, a larger set of external reference objects with well known AB magnitudes will be used to reduce systematic errors. This defines an AB magnitude,

$$m_b^{AB} = m_b^{std} + \Delta_{b-r} + \Delta_r \quad (35)$$

which can be compared to absolute physical flux scales.

**This text needs to be edited into the above section** This self-calibration procedure can be successful only if patches overlap on the sky, so that the same star is observed on multiple patches. This means complete sky coverage is necessary to link all stars together into a rigid system, but also indicates that some amount of dither is required. These investigations have shown that dither patterns where the overlap is one quarter of the field of view or more produce results meeting the SRD requirements.

Note that  $m_{b,i}^{best}$  and  $\delta z_j$  are constrained only up to an arbitrary additive constant. For convenience, this constant can be set so that stars have roughly correct AB magnitudes, however the goal after self-calibration is primarily to have a rigid, self-consistent magnitude system, equivalent to the natural magnitudes. Accurately calibrating the internal magnitudes to an external scale is discussed in the next section, Section 6.

#### 5.4.1. Photometric vs. Non-photometric data

Correcting for the color terms resulting from the difference between  $\phi_b^{meas}(\lambda, t)$  and  $\phi_b^{std}(\lambda)$  requires some preliminary measurement of the color of each calibration star (to within 0.02 magnitudes). This means we must either have some prior knowledge of the colors of each star (from Gaia, for example) or we must have some other method for measuring colors in the  $ugr$  bands relevant to determining metallicity and the color corrections detailed in Section 4.1, presumably by measuring the magnitudes of these stars in photometric data. Without this requirement, we could just combine all photometric and non-photometric data in the self-calibration routine, leaving the self-calibration solver to determine the appropriate  $\delta z_j$  to compensate for any non-photometric images.

Assuming that we must first identify and use photometric data to determine the colors of each object, this could proceed as follows. Identify all observations which were obtained in relatively photometric conditions by searching for images where the average scatter in magnitude measured for each source was less than some threshold (say  $< 0.05$  magnitudes). Using these images and standard stars in these images, measure a preliminary color for each object.

With this preliminary color, make a correction for  $\Delta m_b^{meas}$  and run the self-calibration solver for this (photometric) subset of the data. Iterate the results of the self-calibration solver to improve the color determination for each star, until the color measurement converges to within 0.02 magnitudes.

At this point, we have colors accurate enough to apply a  $\Delta m_b^{meas}$  correction sufficient to run self-calibration on all images, including the non-photometric data. There will be some data which is not calibrateable, due to a large amount of cloud extinction; these images will be identifiable by the low **END**

### 5.5. Calibration Operations

The sequence for photometric calibration is then:

1. Acquire a broadband flat in each filter at the start and end of each observing night. Generate a full, wavelength-dependent illumination correction for the flats on a much longer time interval (timescale to be determined, but much longer than monthly). Apply the appropriate illumination correction to the broadband flat. Apply flat field to images directly.
2. After remaining image processing (bias correction, fringe correction, etc) extract ADU counts of sources from images.
3. Acquire the data cube of narrow-band flat field images, approximately monthly. Apply wavelength-dependent illumination correction. Measure  $\phi_b^{sys}(\lambda, t, x, y)$ .
4. Acquire spectra of known stars roughly every 5–10 minutes throughout each night, fit for atmospheric absorption coefficients and generate  $\phi_b^{atm}(\lambda, t)$  for each science images.
5. Combine  $\phi_b^{atm}$  and  $\phi_b^{sys}$  with a range of model SEDs to create lookup tables for  $\Delta m_b^{meas}$  for various locations in the focal plane.
6. At appropriate intervals (such as at Data Releases), run the self-calibration procedure, applying  $\Delta m_b^{meas}$  to stars chosen for self-calibration procedure and minimizing  $\chi^2$  from equation 32.
7. Apply appropriate  $Z_b^{meas}$  (and potentially  $\Delta m_b^{meas}$  values) to all objects in Data Release catalog, producing standardized magnitudes.
8. Apply measured corrections  $\Delta_{b-r}$  and  $\Delta_r$ , producing absolutely calibrated magnitudes.

This results in calibrated  $m_b^{AB}$  values in a standardized bandpass shape, with above-the-atmosphere fluxes.

## 6. Fixing LSST to an external scale

The next two subsections describe how the internally calibrated standard magnitudes, independently calibrated in each filter bandpass, are fixed to an external scale such that the flux in a single band can be compared to the flux in another filter band (SRD requirement 3) and that the flux in a particular filter band can be compared to an absolute external system (SRD requirement 4). This is equivalent to determining  $\Delta_{b-r}$  and  $\Delta_r$  from Eqn 35. In practice, the same standards will be used for both tasks, and we discuss them together.

Both the band to band calibration for each filter  $b$  (the  $\Delta_{b-r}$  values) and the absolute calibration of  $\Delta_r$  will be determined by measuring the flux from standards whose SEDs across the LSST wavelength range are precisely measured and/or predicted by physical models, and which are contained in the LSST survey fields. Although in principle a single standard with known colors would be sufficient, the major concern here is with systematic errors, for example a variation in color calibration across the sky. Our current simulations (Section ??) show significant spatial patterns in photometric zeropoints. Monitoring and controlling these systematics will require a larger set of standards.

These fall into three groups:

- DA white dwarfs observed with HST. Relatively few in number, these form the "gold standard" for the survey, and include not only the four fundamental standards ((?)), but also a growing set from the HST program of Saha et al ((?)), for which observations are still underway.
- DA white dwarfs for which accurate spectroscopic determination of  $T_{eff}$ ,  $g$ , and interstellar reddening are available.
- Stars observed with Gaia. By roughly 2020, when LSST will be in commissioning, Gaia will achieve 1mmag photometry in the Gaia G band for a large set of stars with  $V < 18.5$  ((?)). If the Gaia G magnitudes can be accurately transformed to the LSST bands, we will have a very large set of absolute standards. This transformation is conceivable, because Gaia produces low resolution spectra for every object. We are currently investigating the photometric accuracy that can be achieved, and do not yet intend to rely on Gaia photometry for calibrating LSST.

Additionally, the main sequence stellar locus can be used to check for systematic errors in photometry.

### 6.1. White Dwarf Standards

Hot hydrogen (DA) and helium (DB) white dwarf stars have simple atmospheres that are well understood (model colors are currently reliable to about 0.01 magnitudes). It is estimated that there will be  $\approx 100/10$  DA/DB WD stars with  $r < 24$  in each LSST image at the South Galactic Pole, though few of these will have available the spectra necessary for synthetic photometry. At least 100–1000 across the sky will be used to search for systematic effects. Catalogs of WD stars visible from Cerro Pachon have been constructed (Bergeron 1992; Bohlin & Gilliland 2004), and a ‘white dwarf calibration system’ has been developed (Holberg & Bergeron 2006).

### 6.2. Population Methods

An alternative approach to using a small handful of precisely known standards is to use the properties of entire populations of astronomical objects.

The locus of main sequence stars in color-color space is also reasonably well understood and has been used to calibrate photometry with success in previous surveys (MacDonald et al. 2004; Ivezić et al. 2007; High et al. 2009). The use of the main sequence stellar locus in addition to WD stars will provide a valuable check on systematic effects that may arise from using (primarily) white dwarfs in the determination of  $\phi^{atm}(\lambda, alt, az, t)$ , as white dwarfs are bluer than most of the main sequence stars used for the bulk of the remainder of the calibration procedures.

Additional checks on the quality of color calibration will be based on color residuals when determining photometric redshifts for galaxies. Analyzing these residuals as a function of galaxy brightness and color, and across the LSST footprint, will yield detailed quantitative estimates of the calibration quality.

### 6.3. Computational Technique for Determining $\Delta_{b-r}$

The values for  $\Delta_{b-r}$  will be determined by generating model  $m_b^{std}$  values for each band-band calibration object, then minimizing

$$\chi^2 = \sum_i \left( \frac{(m_{b,i}^{std} - m_{r,i}^{std})^{meas} - (m_{b,i}^{std} - m_{r,i}^{std})^{model}}{\sigma_{b-r,i}} \right)^2. \quad (36)$$

This comparison can be done using subsets of objects from low Galactic extinction regions, and then bootstrapping to the entire sky to check for systematic effects.

After determining the band to band calibration, there is one further value required to calibrate the entire system to an absolute flux scale:  $\Delta_r$ . This could again be determined using a single object with a well-known flux and spectral energy distribution, however multiple external calibrators provide a valuable check on systematic effects.

Several WDs in the Northern hemisphere have been very precisely calibrated with HST STIS measurements (Bohlin & Gilliland 2004) and it should be possible to obtain similar HST measurements of one or more targets for use in the Southern hemisphere. Identification of these targets has not yet been completed. Nevertheless, as a result of calibration efforts to support the SDSS-II SNe survey, the absolute calibration of the SDSS Stripe 82 region in the  $r$  band is believed to be accurate at the 0.01 mag level (Frieman et al. 2008).

Another route to calibrating to an external flux system is to use standards from Gaia. These will have the advantage of being numerous and widely spread across the sky, with a useful overlapping magnitude range between  $r = 16$  to  $r = 20$ . The magnitude measurements of stars calibrated with the LSST calibration procedure described in this document can be transformed to synthetic magnitudes in Gaia’s bandpasses and compared to Gaia measurements. This comparison across the LSST footprint and as a function of stellar brightness and color could provide a powerful independent test of the quality of LSST photometric calibration, but we are currently unsure of the precision that can be achieved.

## 7. Calibration Hardware

### 7.1. Flat Field Illumination System

The dome screen projector is an array of projectors mounted in the dome of the LSST enclosure, specially designed to provide very uniform illumination to the LSST étendue (the incoming aperture of the telescope, taking into consideration the desired angle of the incoming light) while minimizing any light straying beyond the étendue. In addition, the

dome screen will be able to project both broadband and tunable narrow-band (essentially monochromatic) light sources, providing both broadband (broadband) and narrow-band flat fielding capabilities.

## 7.2. Auxiliary Telescope

reference figure of observing pattern from Burke 2010

Generally, the auxiliary telescope will *not* observe stars along the same line of sight as LSST, as the values for  $C_{mol/O_3/H_2O}$  and  $\tau^{aerosol}$  are better constrained by observing a wide variety of airmasses and locations on the sky that cover a wide range in N/S/E/W directions, as well as utilizing repeat observations of the same star throughout each night, and then fitting the spectroscopic data from the entire night. This improves the signal to noise for the atmospheric absorption profiles generated for each science observation.

## 7.3. Water Vapor Monitoring System

## 7.4. Camera System Telemetry

# 8. Calibration Error Budget

Many sources of errors feed into the calibration process. It is a challenging task to identify and quantify all of them, and we have not yet completely done so. We present our current estimates here, but they are subject to change as our analysis improves. We first discuss repeatability errors, and then uniformity errors. The band-to-band, and absolute calibration errors flow directly from the characteristics of the standard stars, and have already been discussed in Section 6.

## 8.1. Repeatability Errors

The repeatability error in  $m_b^{std}$  for individual calibration stars must meet the requirement in the SRD (Requirement 1 in Section 2). The contributions to this error are best discussed in the context of Figure 4. If we recall that

$$m_b^{std} = m_b^{inst} - \Delta m_b^{obs} + Z_b^{obs} \quad (37)$$

it is natural to divide the error sources into three categories:

- errors in  $m_b^{inst}$ , determined by Flat SED Calibration
- errors in  $\Delta m_b^{obs}$ , determined by SED Correction
- errors in  $Z_b^{obs}$ , determined by Self Calibration

We will assume that the errors in these terms are uncorrelated, and add in quadrature to produce the overall error.

#### 8.1.1. Errors in $m_b^{inst}$

As sketched in Figure 4,  $m_b^{inst}$  is produced by the LSST Data Management System (DM), through processing raw science exposures. Leading error contributions come from the flat spectrum flat (modulated by position and PSF changes of the star on the focal plane), sky background subtraction, and the photometry algorithm used to extract source counts from the flattened and sky-subtracted image. We do not attempt to quantify those individual contributions here, in good part because the relevant software is still early in its development process within DM. Rather, we note that photometric errors, considered as a function of stellar magnitude, always display a systematics floor at the bright end. The contribution of this error to repeatability can not be reduced by any part of the subsequent calibration process, and adds in quadrature to the repeatability error estimate. The estimate from DM is that this error floor will be 3mmag.

#### 8.1.2. Errors in $\Delta m_b^{obs}$

As shown in Section 5,  $\Delta m_b^{obs}$  is a function both of the system bandpass,  $\phi_b^{obs}$ , and the SED of the object,  $F_\nu(\lambda, t)$ . It is affected by the following errors:

- Errors in the atmospheric bandpass,  $S^{atm}$ 
    - Errors induced by measurement noise in atmospheric probe instruments (auxiliary telescope; water vapor monitoring system)
- Measurement errors from the water vapor monitoring system are expected to be +/-1mm. Making use of the plots in Figure 12, which is for a variation of 5mm, this translates to errors of 0.2 mmag in u-band, negligible in g,r, and i-bands, 1 mmag in z, and 2mmag in y.

We do not yet have reliable error levels for aerosols and ozone. Based on the data from (Burke et al. 2010), we estimate that the errors are no more than 10% of the total variation seen in each of those components, RSSed together. Figures ?? and ?? show that the resulting errors are roughly 1.6 mmag in u-, 2.1 in g-band, 0.1 in r, and negligible in all other bands.

- Unmodelled spatial and/or temporal variation of atmospheric components, particularly aerosols

The water vapor extinction is measured along the LSST’s line of sight by the co-pointed microwave radiometer, so it is not subject to this error. For aerosols we again take 10% of the total variation seen. this gives 1.4 mmag in u, 0.7 mmag in g, and negligible for others.

- Errors in determination of monochromatic illumination correction, which propagate directly into errors in  $S_b^{sys}$ .

If the monochromatic illumination correction was ignored, the resulting errors in  $\Delta m_b^{obs}$  are as shown in Figure ???. The maximum effect is roughly 15mmag in u-band and 3mmag or less in the other bands. As discussed in Section 5.2.1, we will determine the illumination correction from a combination of modelling and dedicated measurements. While the accuracy of the result is difficult to assess in advance, we conservatively assume that the error will be 20%. This results in an error contribution of 3mmag in u-band, and 0.6mmag in the others.

- Wavelength-dependent errors in measurement of the domescreen intensity by the photodiode monitor. This will result in errors in combining the monochromatic domeflats to determine  $S_b^{sys}$ .

A photodiode monitor for this purpose was first employed by Stubbs on the PS-1 telescope ((Stubbs et al. 2010)), and an error analysis was undertaken. We are sensitive to systematic errors in the photodiode response rather than to noise, which can be effectively averaged over. Figure ?? estimates the level of these systematic errors as 2 mmag at 400nm, 1 mmag between 470 and 950nm, and increasing beyond 950nm to 10 mmag. Figure 39 shows the effects on Kurucz SEDs of a randomly chosen systematic error curve that conforms to these levels. The effects are negligible in all but the u-band.

- Errors in determining the SEDs of calibration stars from their multicolor photometry.

In practice,  $\Delta m_b^{obs}$  for a calibration star will be determined by looking up its SED as a function of a set of colors formed from the 6 band magnitudes, and then integrating against the system bandpasses  $\phi_b^{obs}$  and  $\phi_b^{std}$  according to equation ???. The SEDs will

come from some model of main sequence (MS) stars (we have been using Kurucz in our work). Real stars deviate from the canonical main sequence (MS) locus, however, due to many causes. Intrinsic widths of the MS color-color locus are generally estimated to be on the order of 20mmag ((??)). This translates directly into an error in  $\Delta m_b^{obs}$ . For example, if we use Figure ?? as an indication of the maximal expected values of  $\Delta m_b^{obs}$ , we can multiply the maximum slope in each band by 20mmag to get an estimate of the effect. The values of the maximum slopes are roughly: 0.28 (u), 0.04 (g), 0.02 (r), 0.02 (i), 0.02 (z), 0.01 (y). This leads to the values in Table 3. This error source is negligible for all but the u-band, where it contributes significantly to the error budget. It is likely that a more careful choice of SED models can reduce this error term.

- Effects of contamination buildup between monochromatic domeflats

Still TBR. This will likely determine the required frequency of taking monochromatic domeflats

- Effects of uncertainties in focal plane temperature on wavelength-dependent detector QE

The detector QE at wavelengths near the red cutoff is affected by temperature, with higher temperatures resulting in higher QE. The focal plane temperature is monitored by the camera, allowing the temperature at any point on the focal plane to be predicted to 0.5 deg K with respect to the reference condition when the monochromatic dome flats were obtained. Figure 40 shows, as expected, that the effect of that prediction error on the natural magnitudes is negligible for all but the y-band, where it is 0.2 mmag. There is, of course, a larger effect on the zeropoints, discussed below.

#### 8.1.3. Errors in $Z_b^{obs}$

Every exposure has a spatial zeropoint model determined by Self Calibration as the solution to a large linear least squares problem, as discussed in Section 5. The errors in the zero points are dependent on the spatial scale of the perturbations, as one can see from a simple argument. One can reasonably expect that Self Calibration will determine the true magnitudes of the calibration stars with an error that is well below the repeatability requirement ( $\sigma \approx 5$  mmag) for individual measurements, so for present purposes we ignore the errors in the true magnitudes. Suppose a perturbation shifts, for a single exposure, the zeropoint over an area  $A$  by some constant  $\delta Z$ . If that area contains  $N$  calibration stars, each provides an estimate of  $\delta Z$  with an error of  $\sigma$ . The full set of calibration stars, assuming that their measurement errors are uncorrelated, allows  $\delta Z$  to be estimated to a precision of

$\sigma_{\delta Z} = \sigma / \sqrt{N}$ . The density of calibration stars varies over the sky. Here we will use a mean value of  $2 \text{ arcmin}^{-2}$ , which corresponds to stars in the range  $20 > V > 18$  at galactic latitude of 30 degrees. This provides the following rough estimates for spatial scales of relevance:

- Detector segment:  $A = 11 \text{ arcmin}^2$ ,  $N \approx 22$ ,  $\sigma_{\delta Z} \approx 1 \text{ mmag}$
- Detector:  $A = 178 \text{ arcmin}^2$ ,  $N \approx 356$ ,  $\sigma_{\delta Z} \approx 0.25 \text{ mmag}$
- Focalplane:  $A = 3.4 * 10^4 \text{ arcmin}^2$ ,  $N \approx 6.8 * 10^4$ ,  $\sigma_{\delta Z} \approx 0.02 \text{ mmag}$

This makes it clear that the size of a detector segment is a rough boundary, above which zeropoint errors should be negligible, while below it they may become quite significant. In reality, many zeropoint perturbations (such as clouds and focal plane temperature variations) have a continuous spatial power spectrum, and a more sophisticated estimate is required.

One can place this argument on a firmer mathematical foundation through "objective analysis", a formulation developed by the oceanographic and meteorological communities, and based on the Gauss-Markov theorem ((?), (?)). The problem addressed by objective analysis is: Given a continuous scalar field, measured in a two dimensional domain at a set of points, what is the optimal estimator for values of the field throughout the domain, and what is the expected error of that estimator? In our application, the scalar field is the zeropoint, and there are (noisy) measurements of it at the position of each calibration star in the field.

We have determined the attenuation factor from simulations of the Self Calibration process (Section 9). Additionally, Gauss-Markov analysis (show results?) **EXPAND**

Zeropoint errors depend not only on the spatial scale of the zeropoint perturbations, as argued above, but also on the magnitude of those perturbations. This is shown in Figures 35, 36 and 37, which present the results of a self calibration simulation which includes *only* zeropoint perturbations, without color-dependent effects. As the first of these figures illustrates, the zeropoint error has the shape of a fan in (zp, zp-error) space, with the width of the fan increasing as the input zeropoint increases. The other two figures, which are cross sections of the fan at different zeropoint levels, show that the error distribution is roughly a gaussian core with fat tails. When the zeropoint is between 0 and 0.5 mag, the gaussian sigma is only 1 mmag. Under extinction conditions near the limit at which we propose to continue survey operations, between 1 and 1.5 mag, the gaussian core has a sigma of 30 mmag. The impact of zeropoint variations on calibration performance as a whole will depend strongly on the pdf of the cloud extinction, as these results show. The strength of this dependence derives partly from the fact that the variance in cloud extinction is roughly proportional

to the extinction itself. Additionally, increasing extinction increases the photometric errors in the measurements of the calibration stars. Our current data on sky conditions at Cerro Pachon supports only a crude measure of this pdf, and there are consequently uncertainties in the overall distribution of repeatability errors.

Perturbations to the zeropoints arise from:

- Clouds, and the varying gray averages of extinction from aerosols, water vapor, and ozone
- Camera gain variation
- Shutter travel time variation
- Detector QE variation due to varying focal plane temperature

## 8.2. Uniformity Errors

The uniformity errors must meet the uniformity requirement in the SRD (Requirement 2 in Section 2). These are driven by systematic errors in  $Z_b^{obs}$  and  $\Delta m_b^{obs}$ . There are several sources:

- errors in standards
- systematic variation of calibration star properties across the sky
- unmodeled systematic variations in the atmosphere

## 9. Testing and Verification

The calibration process presented above is complex, and meeting the SRD requirements is dependent on understanding and controlling a large number of small perturbations. We seek to verify that our approach will produce the required results, and are doing so with different techniques that apply to the stages of LSST’s construction and operation. During the current final design phase, we are employing simulation tools, backed up when possible by measurements on the sky from existing telescopes. (?) and (Burke et al. 2010) are examples of this approach. During the construction phase, we will be able to feed measured data from actual telescope components into the simulations. During operations, our focus is on designing metrics that will let us assess the calibration quality on an ongoing basis.

### 9.1. Self Calibration Simulation

One of the final steps in the LSST calibration process is running self-calibration, solving for a large number of stellar magnitudes and observation zeropoints. The accuracy and precision of the self-calibration procedure has complicated dependencies on various sources of noise as well as the overall observing strategy which determines how well different regions of sky can be tied together. We have undertaken a series of simulations to test the validity of the LSST self-calibration process.

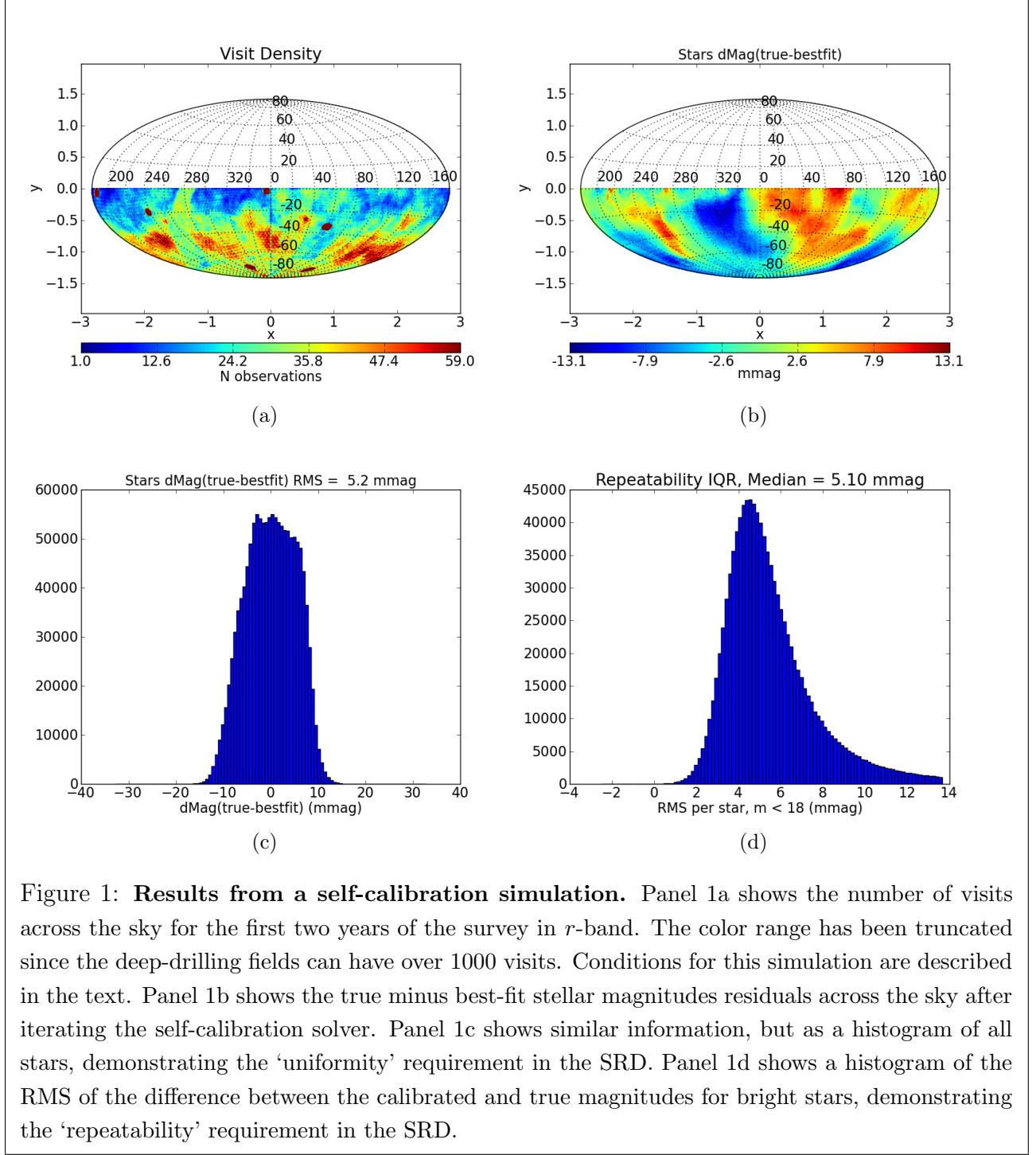
To simulate self-calibration, we require a catalog of realistic LSST observations of bright stars. We generate this catalog by combining Galfast and Opsim. The Galfast code generates a realistic model of the Milky Way stellar populations, matching the observed distributions from SDSS. The LSST Opsim can then be used to generate observations of the stars in the Galfast model. For each Opsim pointing, we generate a list of observed stars along with an observed magnitude that includes atmospheric and hardware effects.

An example self-calibration simulation is shown in Figure 1. This simulates the first two years of LSST observations in  $r$ -band, with 1.3 million stars distributed fairly uniformly across the sky. The observed stellar magnitudes include offsets for cloud extinction, noise based on the Opsim  $5\sigma$ -limiting depth, errors due to ghosting and illumination corrections, and errors based on variations in the filter throughput and placement. The results are promising, showing we easily meet the SRD uniformity requirement, and are very close to meeting the repeatability requirement. The detailed input parameters of this simulation are listed in §C.

Self-calibration simulations will continue to be important during commissioning and regular survey operations. By running simulations that match the conditions of the actual LSST observations, we will be able to identify regions of the sky which are poorly linked to the rest of the survey. We will also be able to identify which spatial scales have the largest errors, which can have important ramifications for cosmological observations (e.g., galaxy clustering) (?).

#### 9.1.1. *Self-Calibration of a Large System Using HEALpixes*

The self-calibration problem for LSST is incredibly large. For comparison, the uber-cal of SDSS involved 36 million observations of 12 million unique stars. For LSST in the first two years we will have around 3.2 billion observations of 100 million unique stars. The memory and computation power required to solve such a system demands that we parallelize the problem in some way.



We have developed a technique using the Hierarchical Equal Area isoLatitude Pixelization (HEALpix) tessellation of a sphere. The HEALpix tessellation was originally designed for analyzing all-sky CMB observations. HEALpixels have equal area, and are distributed to make fast calculations of multipole moments and power spectra.

To run self-calibration in parallel, we assign each observation (consisting of a patch ID, star ID, observed magnitude, magnitude uncertainty, and possibly Illumination patch ID) to the nearest four HEALpixels on the sky. This divides the observations in such a way that we have regions on the sky that are entire HEALpixels plus an added border of approximately one-half HEALpixel. We have had good results with using 768 (53 square degrees) or 3072 (13 square degree) HEALpixels.

Once the observations have been divided, we run the self-calibration solver on each HEALpix region independently. By solving each HEALpixel in isolation, each result has a unique floating zeropoint. To tie the system back to a single floating zeropoint, we construct a matrix based on

$$P_{ij}^{model} = P_i^{best} + HP_j \quad (38)$$

where the  $P_{ij}^{model}$  are the patch zeropoints that were fitted on each HEALpixel. The equation is solved for the true patch zeropoint  $P_i^{best}$  and the HEALpixel floating zeropoints  $HP_j$ , completely analogous to Eqn 33. If an illumination correction is included, we replace the patch ID with a unique ID for each patch ID and illumination ID combination present. At full density, we expect  $\sim 100$  stars per calibration patch and  $\sim 15$  observations per year. Thus, by using the patches to tie large scale solutions together, we have reduced the computational requirements by three orders of magnitude. For example, the two-year  $r$ -band simulation presented shown in Figure 1 has 129 billion non-zero matrix elements in the full self-calibration formulation, but the patch solution needs only 32 million non-zero elements.

This procedure is inefficient in the sense that each patch zeropoint and stellar magnitude is solved for 4 times. However, there is substantial speedup provided by not having to run the solver to convergence over very large spatial scales. Once the final patch zeropoints are solved for, we loop back through the data and apply the zeropoints (and possibly illumination solution) to the stellar observations and calculate a weighted mean for the final best-fit stellar magnitudes. We are currently weighting the patches in Eqn 38 by the number of stars they contain. This should probably be refined, as patches taken in cloudy conditions will be poorly fit even if they contain many stars.

The solutions returned by fitting in parallel are well-matched to solutions which solve the system simultaneously. Figure 3 compares a fit with the traditional global solver with a solution made with HEALpixels.

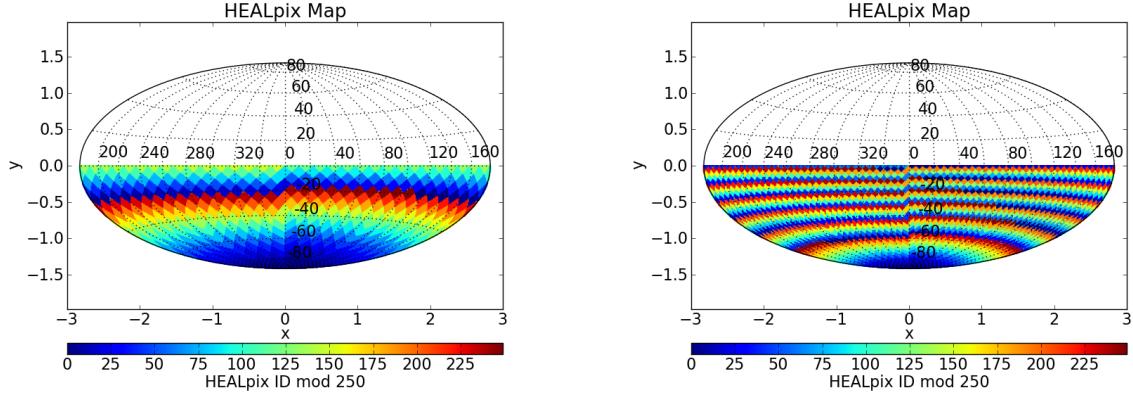


Figure 2: HEALpixel maps for 8 and 16 sides, resulting in 400 and 1568 individual pixels in the Southern hemisphere respectively.

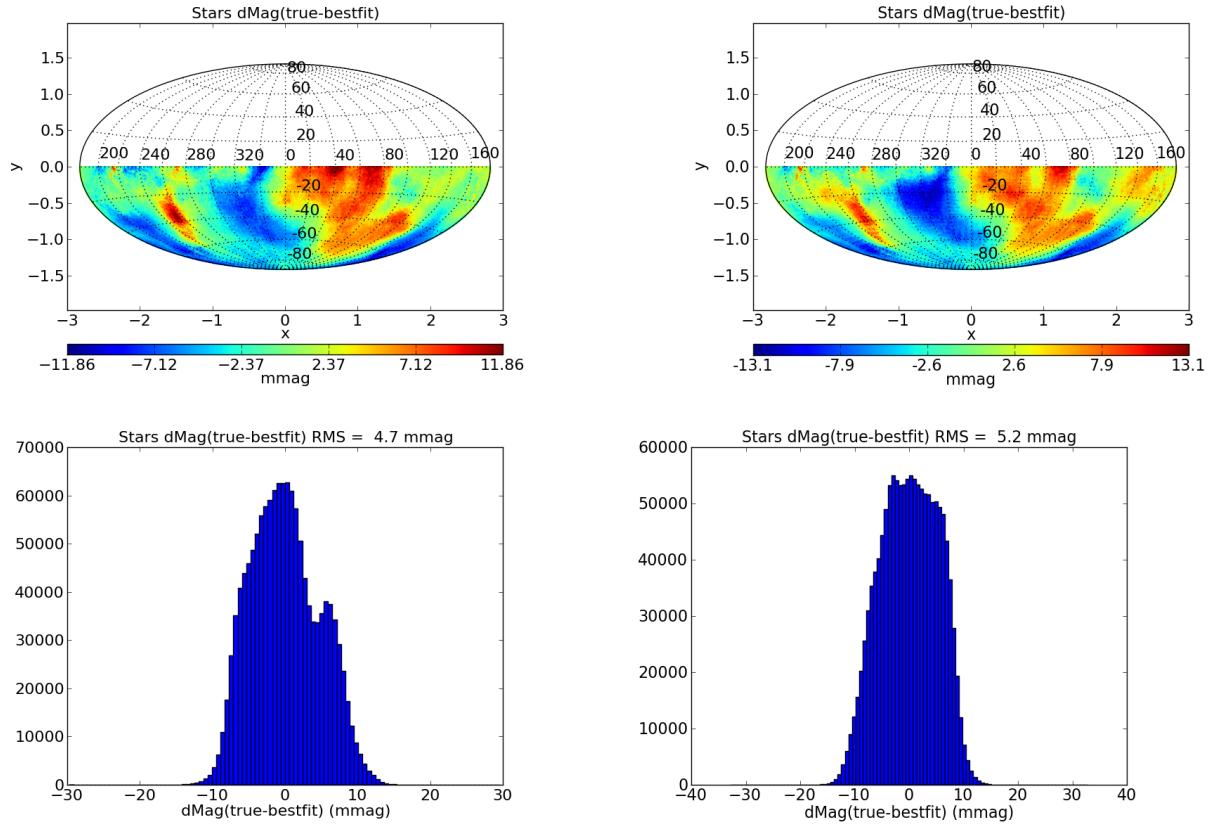


Figure 3: Comparison of a simulation solved simultaneously (left) with one solved on individual HEALpixels and combined. Statistically, the differences between the two solutions are at the sub-millimag level, which is much less than the differences we see between runs with different starting seeds (Figure ??).

## 9.2. Auxiliary Telescope Simulation

### 9.3. Calibration Performance Metrics

#### 9.3.1. Repeatability

XXX-Testing our repeatability is rather trivial, as that does not require any external data. We simply apply the best-fit patch zeropoints to each patch and measure how well we make repeat measurements of each star.

#### 9.3.2. Spatial Uniformity

XXX-We can compare to previous surveys. Pan-Starrs notes that they tend to see SDSS-shaped footprints in their residuals when comparing surveys. This will probably only put an upper limit on the uniformity, as LSST will be deeper and/or have larger coverage than available comparison surveys. Simulations can also help identify regions which we would expect could be poorly linked to the rest of the survey (this is expected to be a problem early in the survey before all regions of the sky have many well-linked observations). We can measure very accurately our RMS as a function of patch and magnitude, and thus generate mock catalog realizations to run through the self-calibration procedure. These mock catalogs should give a good picture of the spatial uniformity.

#### 9.3.3. Flux Calibration

As mentioned in §??, there should soon be a system of white dwarf flux standards with HST observations. We can use subsets of these flux standards to make bootstrap estimates of our overall flux calibration. This is also a potential test of the spatial uniformity of the calibration, although this will be limited if the flux standards are concentrated on the equator.

#### 9.3.4. Color Calibration

There are now a number of techniques for comparing stellar colors across the sky. Ivezić et al. (2004) use a principle color analysis, High et al. (2009) use a stellar locus regression, and Schlafly et al. (2010) look at the color of main-sequence turn-off stars. For all of these techniques, the signal is usually dominated by dust extinction. However, at high galactic

latitudes, the differences in stellar colors reveals errors in the calibration.

In addition to main-sequence stars, our flux standards should be useful for checking the color calibration. As with the flux calibration in a single band, we can exclude some of the flux standards from the analysis and use the excluded stars to see how well we recover their colors.

## 10. Software Implementation

### 10.1. Calibration Products Production

### 10.2. Calibration Within the Data Release Production

### 10.3. Level 2 Data Products

LSST will characterize the system throughput to a degree necessary to achieve the SRD calibration requirements. The system throughput information will be captured by the *normalized system response function*,  $\phi_b$ :

$$\phi_b(\lambda|\mathbf{p}) = \frac{\lambda^{-1} S_b(\lambda|\mathbf{p})}{\int \lambda^{-1} S_b(\lambda|\mathbf{p}) d\lambda} \quad (39)$$

which relates the object's specific flux (spectral energy distribution; SED),  $F_\nu(\lambda)$ , and the calibrated, in-band, flux at the top of the atmosphere,  $F_b$ :

$$F_b = \int F_\nu(\lambda) \phi_b(\lambda|\mathbf{p}) d\lambda \quad (40)$$

In the equations above,  $\mathbf{p}$  captures the dependence on time, telescope pointing (*alt*, *az*), and the position (*x*, *y*) at which the source has been imaged in the focal plane, and  $S_b(\lambda|\mathbf{p})$  is the *total system throughput*. For details, see Eqs 4 and 5 in the LSST SRD.

LSST aims to deliver calibrated fluxes,  $F_b$ , in standard bandpass, for all objects and sources in its catalog. We plan to provide the calibrated fluxes for each source/object computed assuming i) flat  $F_\nu(\lambda) = \text{const}$  SED ( $F_b^{flat}$ ), and ii) an SEDs chosen from a library of SEDs, that is consistent with object properties (measured color being the most apparent one, but shape, Galactic longitude and latitude, and others may be considered as well). We expect most users will prefer and use the latter.

These assumed SEDs and precomputed fluxes will not be appropriate for all cases. Examples include objects exhibiting SED variability, and objects with exotic SEDs. To

enable *recalibration* of the measured fluxes using user-supplied SEDs, we will retain and make available the normalized system response function,  $\phi_b$ , for every visit. This will allow the computation of the multiplicative corrections to the provided flux:

$$c = \frac{\int f_{\nu}^{user}(\lambda) \phi_b(\lambda|\mathbf{p}) d\lambda}{\int f_{\nu}^{flat}(\lambda) \phi_b(\lambda|\mathbf{p}) d\lambda} \quad (41)$$

and computation of the recalibrated flux as:

$$F_b^{user} = c F_b^{flat} \quad (42)$$

Note that  $f_{\nu}^{flat}(\lambda) \equiv 1$  and since  $\phi_b$  is normalized to one, the denominator of Eq. 41 is identically equal to one as well. Therefore, only the numerator needs to be computed to get  $c$ .

#### 10.3.1. Storing and Obtaining $\phi_b(\lambda|\mathbf{p})$

The total system throughput,  $S_b(\lambda|\mathbf{p})$ , from which  $\phi_b(\lambda|\mathbf{p})$  is derived, can be factored into two components,  $S_b^{sys}$  and  $S_b^{atm}$ :

$$S_b(\lambda, t, alt, az, x, y) = S_b^{sys}(\lambda, x, y, t) \times S_b^{atm}(\lambda, alt, az, t) \quad (43)$$

, capturing its dependence on the telescope+camera system and the atmosphere, respectively. The two will be measured and/or derived as described in Document XXX.

$S_b^{sys}$  will be stored as a series of  $(\lambda, x, y)$  data cubes. For each measurement, we expect on order of  $\sim 100$  bins in both  $x$  and  $y$ , and about  $\sim 200$  bins in  $\lambda$ . As discussed in Section 5,  $S_b^{sys}$  is expected to vary slowly with time, necessitating not more than  $\sim$ monthly re-characterization (and therefore, about one new data cube per month). These data cubes will be made available to the user, both in bulk form (e.g., FITS files), and as tables (or functions) in the LSST database.

$S_b^{atm}$  will be computed from models of the atmosphere derived using the spectra obtained with the auxiliary telescope. The parameters of this model will be stored in the database and made available to the user, together with the software needed to compute  $S_b^{atm}$ . For ease of use, we will also provide database tables or functions returning  $(\lambda, S_b^{atm}, \sigma_{S_b^{atm}})$  evaluated with  $\Delta\lambda = 1\text{nm}$  resolution, for each visit<sup>2</sup>.

---

<sup>2</sup>Or potentially for each CCD in a visit, in case of strong gradients in  $S_b^{atm}$

The radiometer/GPS data on precipitable water vapor will be stored in the database together with other exposure metadata. The user will be able to query this information for each exposure.

#### 10.3.2. Database-level Recalibration

We will provide an easy-to-use database interface to perform recalibration given user-defined SEDs. While the exact syntax is yet to be finalized, the following example should be illustrative of how the user could obtain recalibrated time series using the highest-level of provided APIs:

```
SELECT
    midPointTai,
    filter,
    recalibPsFlux(sourceId, mySeds.sedId, "mySeds")      as flux,
    recalibPsFluxErr(sourceId, mySeds.sedId, "mySeds") as fluxErr,
FROM
    Sources, mySeds
WHERE
    Sources.objectId = mySeds.sedId
INTO
    myRecalibratedMagnitudes
```

The query above assumes that, for each object of interest, the user has uploaded their SED into a table named `mySeds` and keyed it by `objectId`. Given the name of the table and the ID of the SED entry, functions `recalibPsFlux` and `recalibPsFluxErr` will use the `sourceId` to locate the appropriate  $S_b^{sys}$  and  $S_b^{atm}$  pertaining to that specific measurement, compute  $\phi_b$  and the correction  $c$ , and return the recalibrated point source flux and its error, respectively.

A suite of lower level functions will be provided as well, providing a more granular control of the recalibration process.

## 11. Risks and Mitigations

The error budget presented in Section 8, and the simulation results presented in Section 9 show that LSST will meet SRD requirements for calibration, but with small margin.

There is clearly a risk that once commissioning or survey operations begin, an unanticipated source of error will arise, or that known error sources will have been underestimated. The calibration process that we have designed is very similar to those of PS-1 and DES, as previously mentioned, though with some significant improvements. It is therefore cause for some concern that the PS-1 results reported in (Tonry et al. 2012) are significantly poorer than we would expect from our error analysis, and show clear evidence for systematic errors that are not understood. Before turning to more specific risks, we first address this general concern.

One mitigation that we will explore involves a modification of the overall calibration process. As presented in Sections x and y, the current design cleanly separates compensation for changing system bandpass shape (performed in SED Correction 4), and for changing zeropoints (performed in Self Calibration). The SED Correction process is based completely on forward modeling, in which the effect of some wavelength-dependent perturbation (eg change in atmospheric aerosols) on a particular calibration star is predicted by a measurement-based model of the effect. As an alternative, one can imagine combining the SED Correction and Self Calibration processes into one large least squares problem, in which not only the zero-points are determined, but all the wavelength-dependent effects as well.

We have tested a first step in this direction by incorporating a parameterized model of the wavelength-dependent illumination into Self Calibration. Initial results have not been promising, but this is far from invalidating the overall approach.

A further risk is clearly that our overall conclusion about the performance of the calibration process is based largely on simulation tools. These tools require significant further development, and completion of a validation process.

- Systematics error floor for photometry is higher than expected
- Forward models for illumination correction are less accurate than expected
- Stability of  $S_b^{sys}$  is less than expected

## REFERENCES

- Bergeron, P. 1992, JRASC, 86, 309  
Bohlin, R. C., & Gilliland, R. L. 2004, AJ, 128, 3053  
Burke, D. L., et al. 2010, ApJ, 720, 811

- Eppeldauer, G. P., Yoon, H. W., Zong, Y., Larason, T. C., Smith, A., & Racz, M. 2009, *Metrologia*, 46, 139
- Frieman, J. A., et al. 2008, *AJ*, 135, 338
- Hansen, O. L., & Caimanque, L. 1975, *PASP*, 87, 935
- High, F. W., Stubbs, C. W., Rest, A., Stalder, B., & Challis, P. 2009, *AJ*, 138, 110
- Holberg, J. B., & Bergeron, P. 2006, *AJ*, 132, 1221
- Ivezić, Ž., et al. 2004, *Astronomische Nachrichten*, 325, 583
- . 2007, *AJ*, 134, 973
- Kurucz, R. L. 1993, *VizieR Online Data Catalog*, 6039, 0
- MacDonald, E. C., et al. 2004, *MNRAS*, 352, 1255
- Marshall, J. L., et al. 2013, *ArXiv e-prints*
- Nugent, P., Kim, A., & Perlmutter, S. 2002, *PASP*, 114, 803
- Padmanabhan, N., et al. 2008, *ApJ*, 674, 1217
- Radomski, J., Trancho, G., Fuhrman, L., Falvey, M., Gigoux, P., Montes, V., Daruich, F., & Lazo, M. 2010, in Society of Photo-Optical Instrumentation Engineers (SPIE) Conference Series, Vol. 7737, Society of Photo-Optical Instrumentation Engineers (SPIE) Conference Series
- Regnault, N., et al. 2009, *A&A*, 506, 999
- Schlafly, E. F., Finkbeiner, D. P., Schlegel, D. J., Jurić, M., Ivezić, Ž., Gibson, R. R., Knapp, G. R., & Weaver, B. A. 2010, *ApJ*, 725, 1175
- Schlafly, E. F., et al. 2012, *ApJ*, 756, 158
- Stubbs, C. W., Doherty, P., Cramer, C., Narayan, G., Brown, Y. J., Lykke, K. R., Woodward, J. T., & Tonry, J. L. 2010, *ApJS*, 191, 376
- Stubbs, C. W., & Tonry, J. L. 2012, *ArXiv e-prints*
- Stubbs, C. W., et al. 2007, *PASP*, 119, 1163
- Tonry, J. L., et al. 2012, *ApJ*, 750, 99

Vanden Berk, D. E., et al. 2001, AJ, 122, 549

Veselovskii, I., Whiteman, D. N., Korenskiy, M., Kolgotin, A., Dubovik, O., & Perez-Ramirez, D. 2013, Atmospheric Measurement Techniques Discussions, 6, 3059

Wittman, D., Ryan, R., & Thorman, P. 2012, MNRAS, 421, 2251

### A. Filter Set

Figure 41 illustrates the baseline LSST filter bandpasses, including a ‘standard’ atmosphere, and baseline estimates for the mirrors, lenses, filter and detector transmission and sensitivity functions.

### B. Photometric measurements for non-main sequence stars

LSST will record a series of  $m_b^{nat}$  measurements for each astronomical object in each visit. These  $m_b^{nat}$  measurements are generated directly from the counts recorded in each image, corrected with the photometrically uniform, broad-band flat field and for gray (cloud) atmospheric extinction effects. However, these  $m_b^{nat}$  measurements will vary as the shape of the bandpass changes, whether as a function of position in the focal plane or as a function of changes in atmospheric absorption components. Correcting for these effects requires assuming a particular SED for each source, and produces  $m_b^{std}$  values after applying  $\Delta m_b^{meas}$  offsets (see the overview of calibration in section 5 for a review).

For most objects, LSST will simply assume the source has a flat SED, which means that  $m_b^{std}$  is exactly equal to  $m_b^{nat}$ , as by definition then  $\Delta m_b^{meas}$  is zero. To permit scientists to generate higher precision photometry for objects using arbitrary SEDs, LSST must provide a record of  $\phi_b^{meas}(\lambda, alt, az, x, y, t)$  as well as the normalization zeropoint offsets for each observation. With these additional pieces of information, scientists can generate more appropriate  $\Delta m_b^{meas}$  corrections, using their own chosen object SED to generate  $m_b^{std}$ . Sections ?? and ?? outline the typical magnitudes of these corrections for main sequence stars;  $\Delta m_b^{meas}$  can easily be on the order of 20 mmag for *gri*, or even 100 mmag in *u* band. For more extreme SEDs, these corrections may be even larger.

Figure 43 illustrates the likely magnitude of these  $\Delta m_b^{meas}$  corrections for a wide variety of SEDs. In each plot, the main sequence stars are shown as in the figures in the main paper (small dots, color-coded by metallicity), although given the increased scale here they only appear as a purple series of circles. M dwarfs are now included, generally mimicking the behavior of the main sequence stars but extending further into the red. More unusual SEDs are also included; a quasar SED, based on a composite of many empirical quasars from SDSS from Vanden Berk et al. (2001) that has been extended to the full LSST wavelength range through the addition of power law flux above and below the original range ( $f_\nu \propto 1/\lambda^{0.5}$  for  $\lambda < 89\text{nm}$  &  $f_\nu \propto 1/\lambda^{1.5}$  for  $\lambda > 800\text{nm}$ ), and redshifted from  $z = 0$  to  $z = 3$ ; also a sample of SN Ia from templates generated by Peter Nugent (Nugent et al. 2002), redshifted from  $z = 0$  to  $z = 1$ .

The figure shows the  $\Delta m_b^{meas}$  values that would be expected under a maximum change of atmospheric parameters and under a likely bandpass shift. This demonstrates how much the reported  $m_b^{nat}$  values could vary for each object. If LSST was to just calculate an offset between  $m_b^{nat}$  and  $m_b^{std}$  based on an object’s color (and assuming that the object had an SED similar to a main sequence star), the resulting  $m_b^{std}$  values would be incorrect by the value of the offset between the true  $\Delta m_b^{meas}$  for the SED and the main sequence  $\Delta m_b^{meas}$  values at each color; this could easily be more than 20mmag.

With the wide variety of objects and SEDs that will be possible in the LSST data set, it is crucial to provide the information and tools so that scientists may calculate highly precise  $m_b^{std}$  magnitudes for their objects of interest. This means recording or providing the means to recalculate  $\phi_b^{sys+atm}(\lambda, t)$  for every object detected with LSST.

### C. Fiducial Self-Calibration Input

The input file used with simSelfCalib.py to generate the simulated LSST catalog presented in §9.1.

```
# the number of total stars in the sim
nStarTot = 2000000
# the magnitude range
magmin = 17
magmax = 21
# the color (g-i) range of the sim stars
colmax = 3.5
colmin = -0.8
# A value for the random seed, if desired.
random_seed = 42
# Distribution to use for mag_rand_err. Options: Gaussian, Cauchy
errorDist = Gaussian
# Parameters for using flux standard stars.
# Flux standards are all placed on patchid 0.
# number of stars per block to use.
# values < 1 make a fraction of the stars standards,
# values > 1 are the number of stars per block.
fluxFrac = 0
# Gaussian noise to add to the flux standard mags.
fluxNoise = 0.0001
```

```
# the next few values are about the zeropoint and color terms
# the random error added to the stars, distribution set by errorDist.
# This is an error floor.
    mag_rand_err = 0.003
# calculate what LSST will think the error is based on observed mag (True) or output a
# more realistic error (False)
    calcerror = False
# Use Opsim xparency for the gray zeropoints
    zp_opsim = True
# the maximum zeropoint change added to the patches
    zp_var_max = 1.0
# the maximum zeropoint gradient added to the patches
    zp_grad_max = 0.0
# the maximum color term change in mag added to stars in a patch
    colterm_max = 0.005
# the range between -1 < (g-i) < 1 in the color term due to filter bandpass shift.
    colterm_rad_max = -0.040
# the uncertainty in the observed color -- 0 or negative values will result
# in no color gradient correction
    color_obs_noise = 0.05
# the fractional uncertainty in the color correction term ()
    color_correction_error = 0.10
# fraction of the FoV that the filter can jitter:
    filter_jitter = 0.0055
# fractional variation in the gain (Gaussian RMS)
    gainvar = 0.00
# fractional error in the exposure time (e.g., from shutter errors)
    exptvar = 0.001
# number of HEALpixel sides to use. Must be power of 2.
    nside = 16
# Use cloud images. Each visit has a cloud image generated on a sparse
# grid then interpolated to the star positions.
    use_cloudsimage = True
# the magnitude of the sinusoidal variation
# in the zero point in the X and Y-directions. Note, these
# interfere with eachother so the total error
# will be \pm (sinvarx_mag+sinvary_mag)
    sinvarx_mag = 0.00
```

```
    sinvary_mag = 0.00
# the spatial scale for the sinusoidal variation
# value of 1 will make the FOV run from -pi to pi, 2--> -2pi to 2pi
    sinvarx_scale = 1.
    sinvary_scale = 1.
    sinx_phase = 0.0
    siny_phase = 0.0
# angle between the x and y variation axes. This is held constant.
    sinx_angle = 20.0
# a second set of spatially varying sinusoidal errors.
# Here, the phase is randomly set for each night to simulate flat fielding errors
    flat_sinvarx_mag = 0.00
    flat_sinvary_mag = 0.00
    flat_sinvarx_scale = 1.
    flat_sinvary_scale = 1.
# angle between x and y axes. Held constant, but the phase is
# varied from night-to-night
    flat_sinx_angle = 0.0
# how many nights to keep the phase constant
    phase_persist = 1
# illumination correction file jdsu_utest.dat
# illumcorr_filename = None
#
# illumination correction file jdsu_r.dat  ricbb.dat
#
    sed_illumcorr_filename = jdsu_r.dat
    bb_illumcorr_filename = ricbb.dat
# what fraction of the illumination error to use
    illum_err_factor = 0.1
# focal plane temperature model file
# fpModel1_filename = fpTemp1.dat
# fpModel2_filename = fpTemp2.dat
# spatial variation from cloud removal
# fraction of the total zeropoint shift to put in as the cloud magnitude
    cloud_mag = 0.00
    cloud_sinvar_scale = 5.
# Variable star contamination
    rr_fraction = 0
```

```
    rr_amplitude = 1.  
# Use Kepler stellar variability statistics?  
    kepler_variability = False  
# limits for sim footprint  
# the ra/dec limits, in degrees  
    raMin = 0  
    raMax = 360.  
    decMin = -90.  
    decMax = 0.  
# you can use opsim (True or False) to generate fields  
# use opsim database (opsim3.61)  
    use_opsim = True  
# flag to use the opsim dither scheme, instead of random dither.  
# note that opsim dither scheme covers the entire  
# radius_fov (i.e. dith_Offset_frac = 1)  
    use_opsimdither = True  
# opsim filter, if using this  
    opsimfilter = r  
# then specify time to start/stop, in nights from start of opsim  
    tstart = 0  
    tstop = 730  
# or don't use opsim and specify nEpoch instead  
# nepoach = number of visits to each field  
    nEpoch = 10  
# you can use calsim database to generate the stars.  
# Available calsimtables include: "msrgb", "msrgb_1e6", and "msrgb_1e7"  
    use_calsim = True  
    calsimtable = msrgb_1e6  
# you can also change the radius of the field of view  
    radius_fov = 1.8  
# and the number of patches to split this radius into. should be 'square' (nPatch=N^2)  
    nPatch = 25  
# raOff and decOff control dithering  
    dith_raOff_frac = 0.5  
    dith_decOff_frac = 0.5  
# Limits of camera rotation, in degrees (has a dithering effect)  
# (LSST standard = -90 to 90)  
    dith_skyRot_min = -90
```

```
dith_skyRot_max = 90
# filenames for outputs
# starobs is the file input to the solver
starobs_filename = star_obs.dat
# should starobs contain the subpatch information?
# These can then be solved for illumination corrections.
print_subpatch = True
# number of radius bins to use for the illumination correction
nRadiusBins = 5
# number of g-i color bins to use for the illumination correction
nColorBins = 1
# output each ra/dec block to it's own file (only use for very large sims)
multifile = False
# master file contains everything relevant to every measurement of every star
master_filename = /dev/null
# stardata is the basic star magnitudes, positions and colors
stardata_filename = stardata.dat
# star data for re-assigning patches
ras_filename = /dev/null
# visitdata gives the information on the visits used in the simulation
visit_filename = visit.dat
# patchdata gives the information on the patches in the simulation
patch_filename = patchdata.dat
```

## D. Glossary

- **Level 1 Data Product.** A data product, such as a measurement of an astronomical object’s position or flux in a single image, that is computed on a nightly basis. Level 1 data products primarily consist of alerts on transient, variable and moving objects. The photometric calibration process outlined in this paper does not apply to Level 1 data products. Level 1 data products will be calibrated using all applicable prior knowledge (including secondary standard catalogs generated from previous Data Release calibration of all LSST-observed stars in the field).
- **Level 2 Data Product.** A data product, such a measurement of an astronomical object’s position or flux in either a single image or a series of images, that is computed on the Data Release schedule, on a six-month or yearly schedule. Level 2 data products

leverage all previous observations of the same object, as well as all knowledge of the LSST system accumulated to that point. The photometric calibration process outlined in this paper is used to generate Level 2 data products.

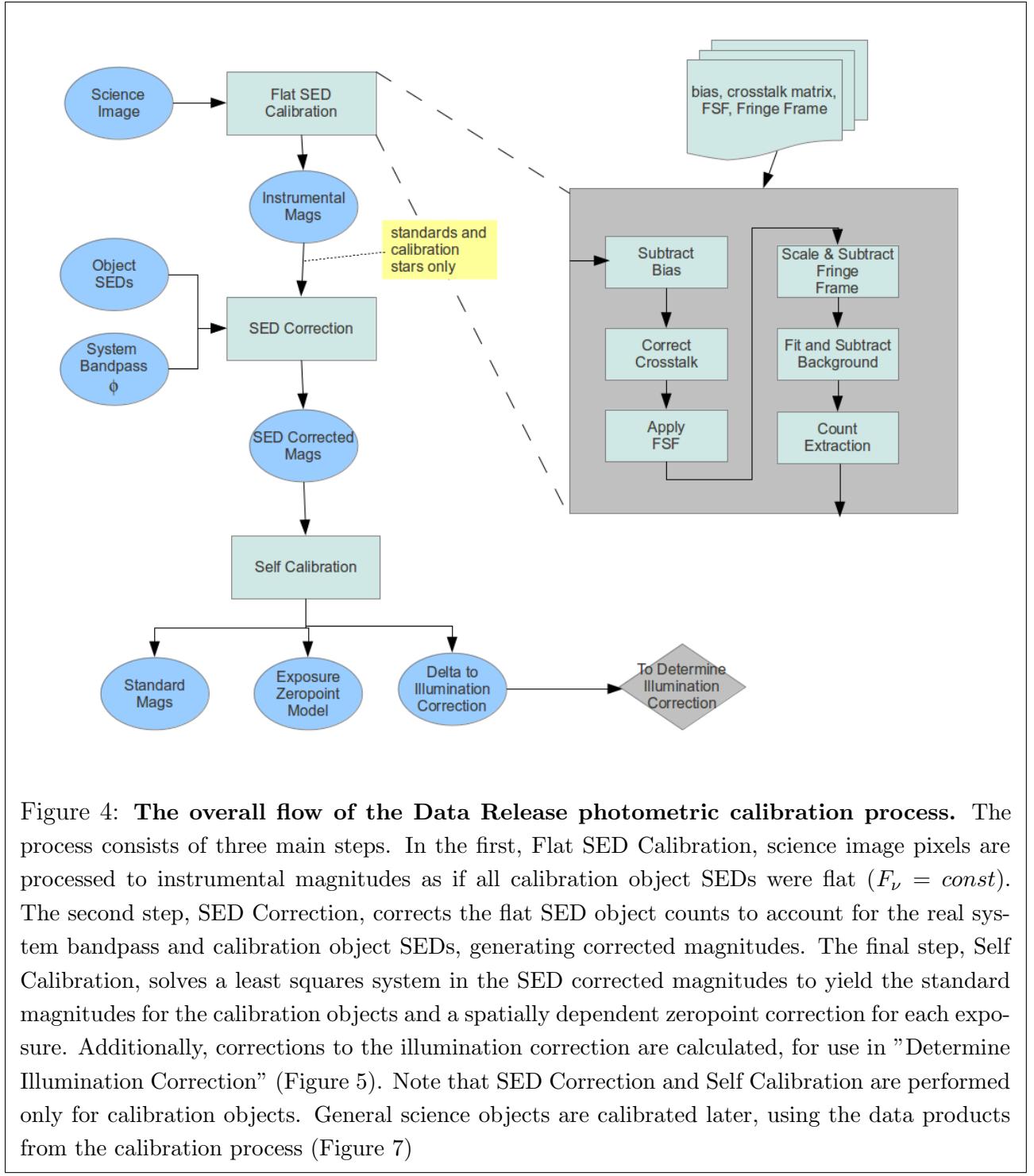
- **Normalized system response,**  $\phi_b(\lambda)$ . The normalized system response describes the shape of the bandpass transmission curve, separating this from the normalization of the throughput curve which can be determined separately.  $\phi_b(\lambda)$  is described by Equation 5. The integral of  $\phi_b(\lambda)$  is always 1.
- **Camera Calibration Optical Bench (CCOB).** The CCOB is an apparatus to calibrate the spatial and wavelength-dependent response of the focal plane (detector + camera). The CCOB uses a well controlled, wavelength-variable, light source (such as a tuneable laser) calibrated using a NIST photodiode to illuminate the focal plane when the camera is unmounted from the telescope. This light source, which produces a spot in the focal plane approximately the size of or smaller than the PSF, will be scanned across the detector  $(x, y)$  at a variety of beam incident angles,  $(\theta, \phi)$  and at a variety of wavelengths  $(\lambda)$ . This allows the response of the detector to be measured in the presence of a well-understood light source. The response of the detector can be measured in two different configurations: one with only the detector and the dewar window - which doubles as lens 3 (L3) - and one with the detector, L3, L2, L1, a small test-section of filter and the camera shutter. The filter test section used is not the full LSST filter, and thus will not capture spatial non-uniformities in the filter bandpass. The CCOB provides test data about the camera assembly for camera acceptance and will help constrain the optical ZEMAX model, although without a full filter it cannot capture the full set of parameter required for the ZEMAX model. More details about the requirements and physical apparatus of the CCOB are available in LSST-10015 and LSST-8217.
- **Broadband flat field.** An image obtained by observing a light source which generates photons with a wide range of wavelengths (as opposed to a narrow-band flat), with relatively uniform illumination across the field of view. Night sky flats, twilight flats, and white-light or broadband dome screen flats would all generate broadband flat fields.
- **Narrow-band flat field.** An image obtained by observing a light source which generates photons with a very narrow range of wavelengths (hypothetically, even a single wavelength), with relatively uniform illumination across the field of view. A dome screen illuminated with a narrow-band laser light source will generate a narrow-band flat field.

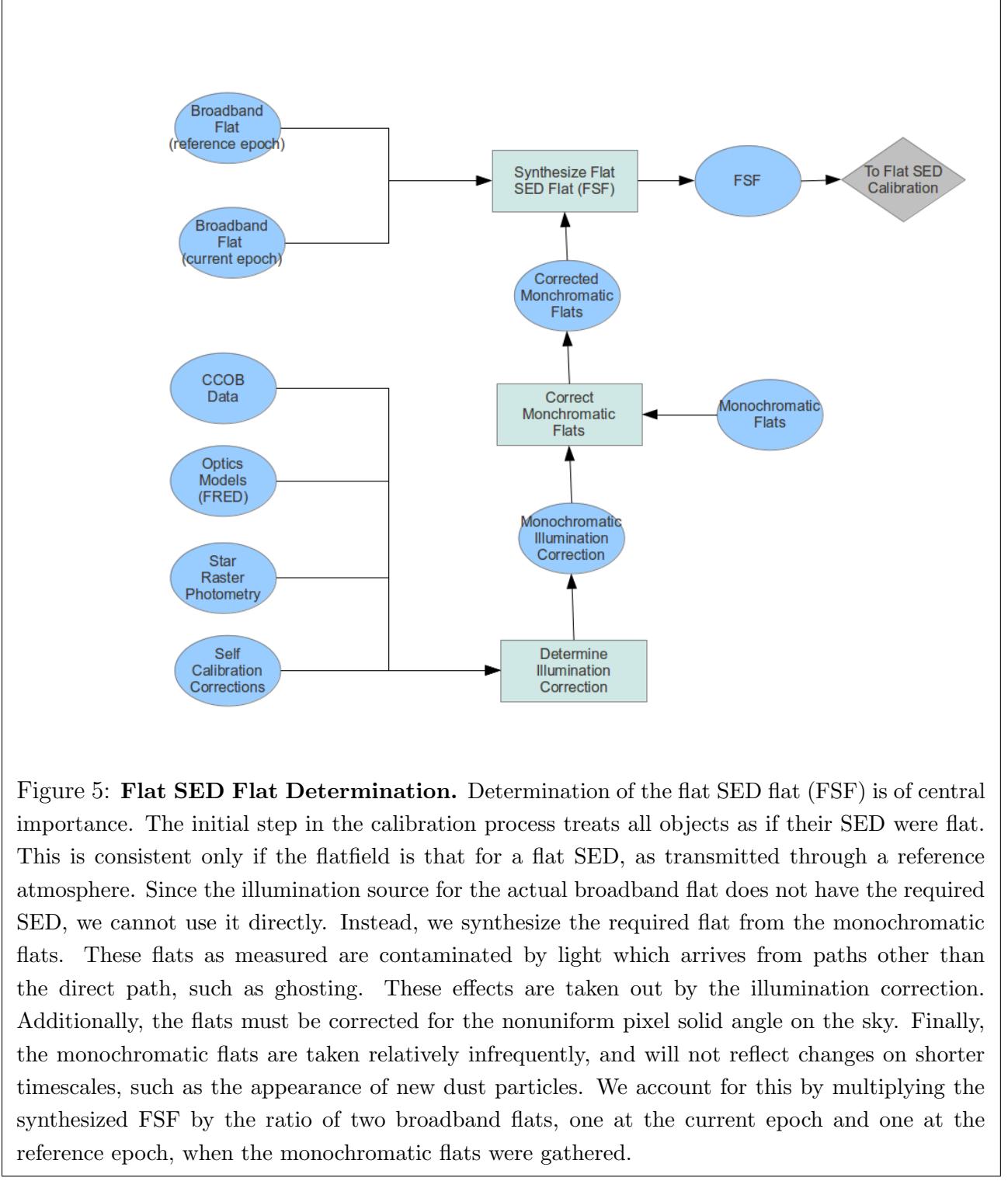
- **‘Observed’ flat field.** A flat field, as obtained by observing the dome screen projectors. Generally would refer to a broadband flat field.
- **Photometric flat field.** A flat field which produces uniform photometric measurements across the field of view for a flat  $F_\nu(\lambda)$  source. A photometric flat field must be based on a broadband flat.
- **Illumination Correction.** The ratio between the photometric flat field and the observed flat field.

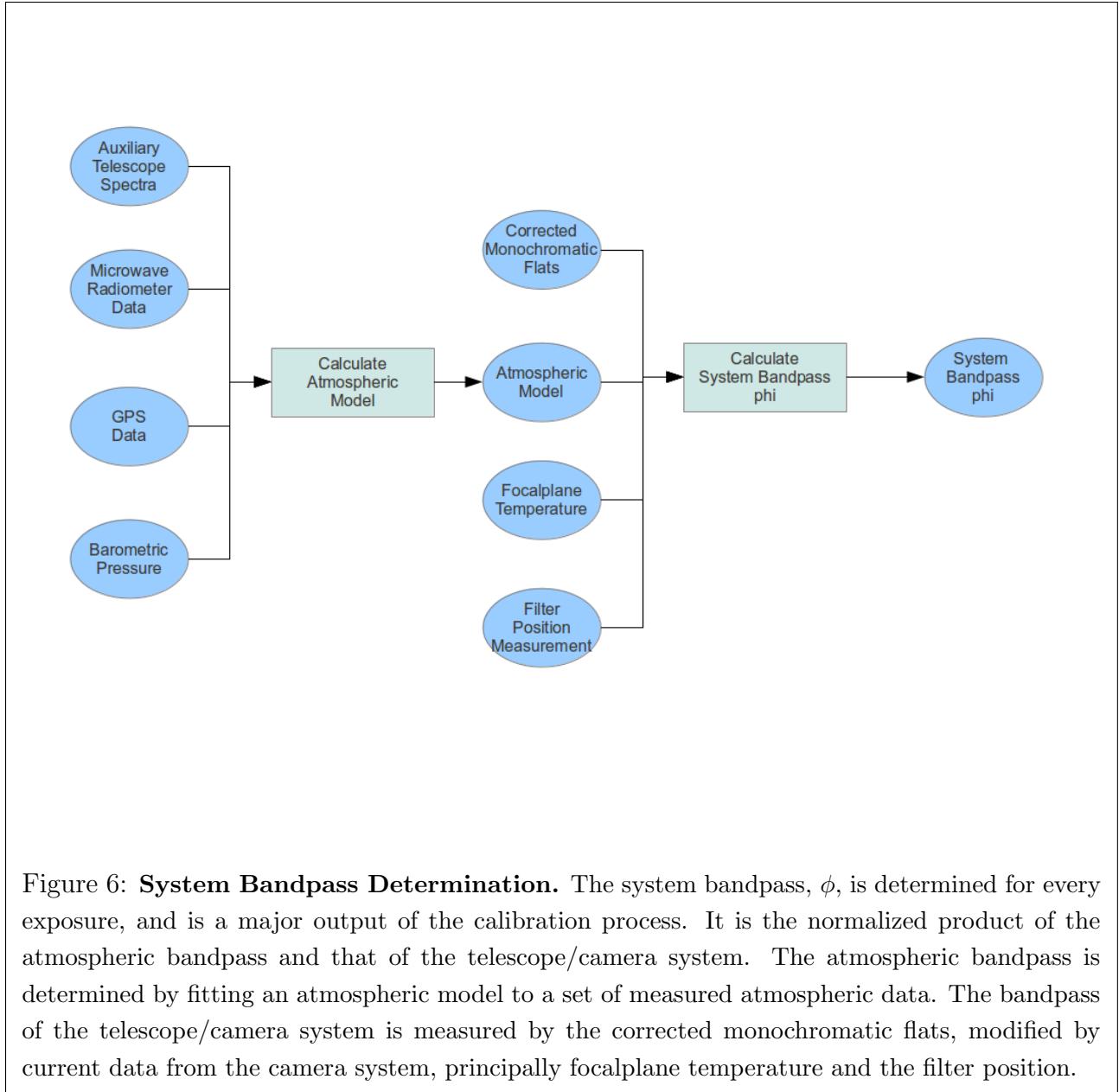
$$\text{Flat}_{\text{photometric}} = \text{Flat}_{\text{observed}} * \text{IlluminationCorrection} \quad (\text{D1})$$

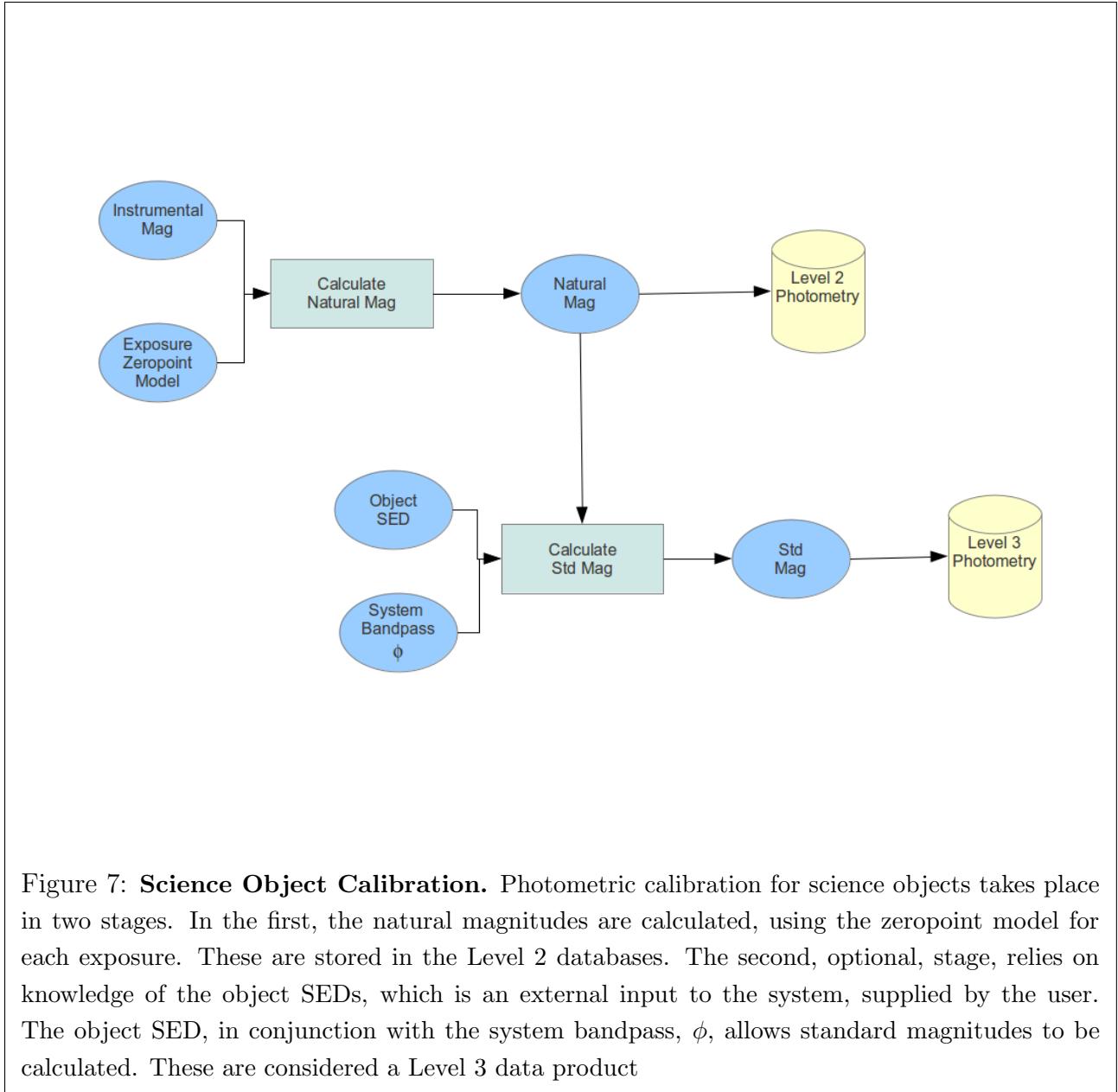
- **Natural magnitude.** A magnitude measurement which relates directly to the number of counts measurement in an image (after including a photometric flat field correction and a rough zeropoint for an entire image). The natural magnitude relate to an ADU count that does *not* account for the color or SED of the source being observed, thus does not include any wavelength-dependent corrections. For a non-variable source observed under variable atmospheric transmission conditions and/or at varying locations in the field of view, the natural magnitude reported will change due to changes in the bandpass shape. The natural magnitude is equivalent to an observed magnitude, after the appropriate zeropoints have been applied.
- **Standard magnitude.** A magnitude measurement which includes not only corrections for the photometric flat field and a rough zeropoint for the image, but also includes a correction for wavelength-dependent effects. This means the  $\Delta m_b^{\text{meas}}$  appropriate to correct the natural magnitude of the object from the observed bandpass shape,  $\phi_b^{\text{meas}}(\lambda, t)$ , to the standard bandpass shape,  $\phi_b^{\text{std}}(\lambda)$ , has been calculated for the SED of the object and applied. For a non-variable source,  $m_b^{\text{std}}$  will be constant over time even if the atmospheric absorption curve or the location in the field of view changes.
- **Operations Simulation.** The Operations Simulation is a simulated pointing history of LSST, covering the sky in the same manner as the telescope could, in practice. It uses weather conditions based on historical records from Cerro Tololo, including appropriate seeing and sky brightness variations. The motion of the telescope is simulated in high fidelity, including acceleration from field to field and cable wrap. A variety of proposals are used to determine which fields to observe at each time; these proposal include the ‘universal cadence’ (satisfying most of LSST’s science requirements) and ‘deep drilling’ (a limited set of fields, observed frequently and deeply over the lifetime of the survey).

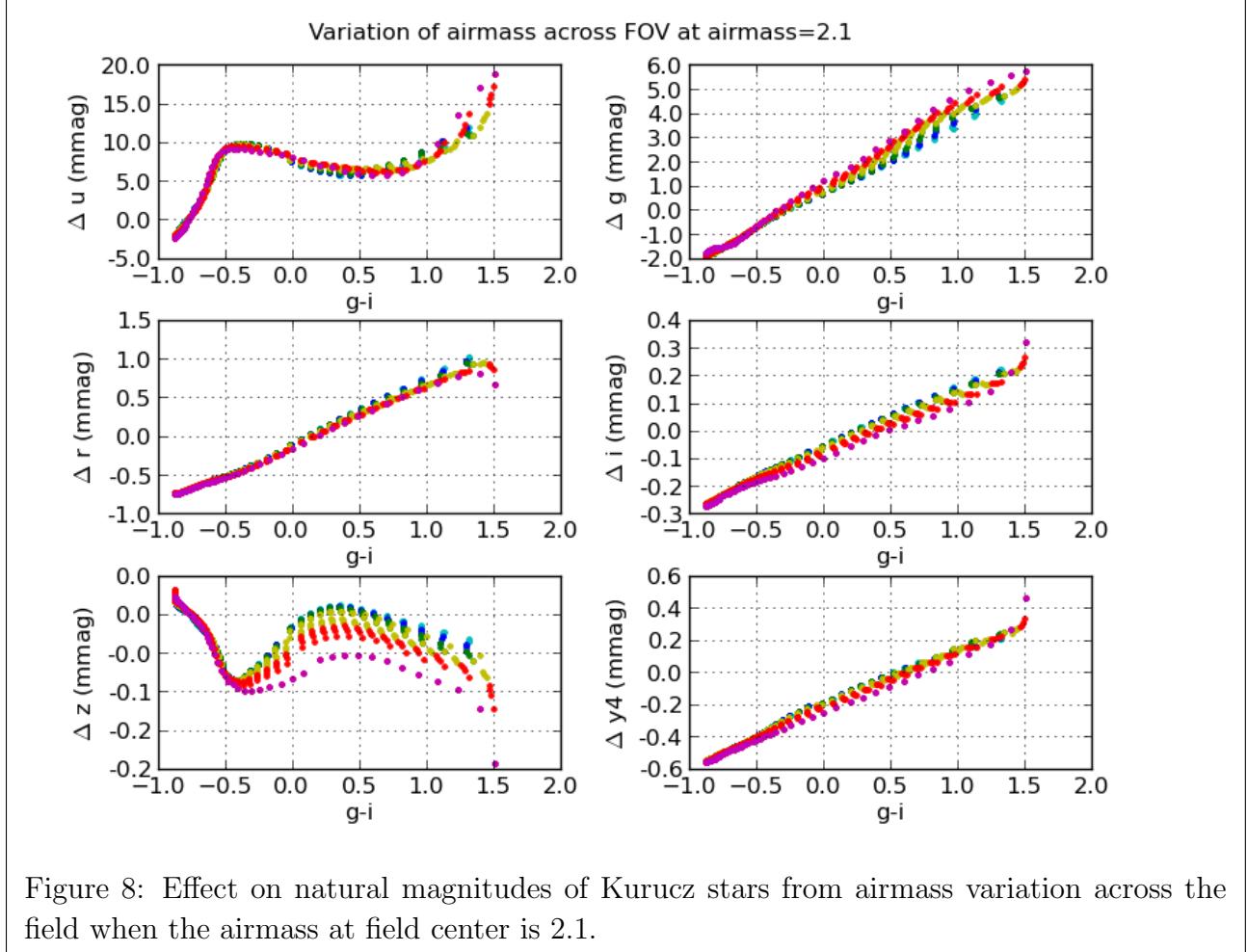
- **PWV.** Precipitable Water Vapor. The total column depth of water vapor in the atmosphere, measured at zenith. The units are mm of liquid water equivalent.

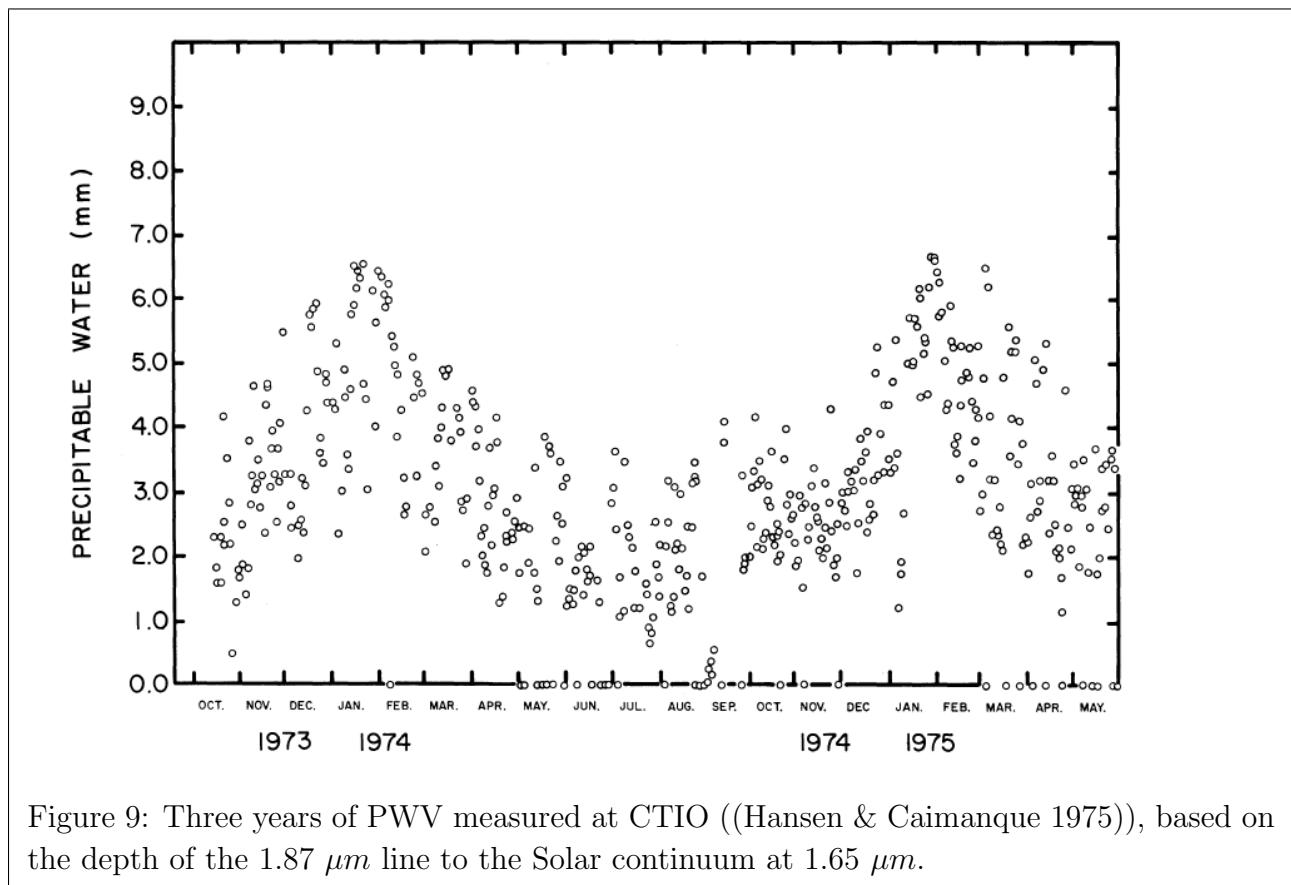


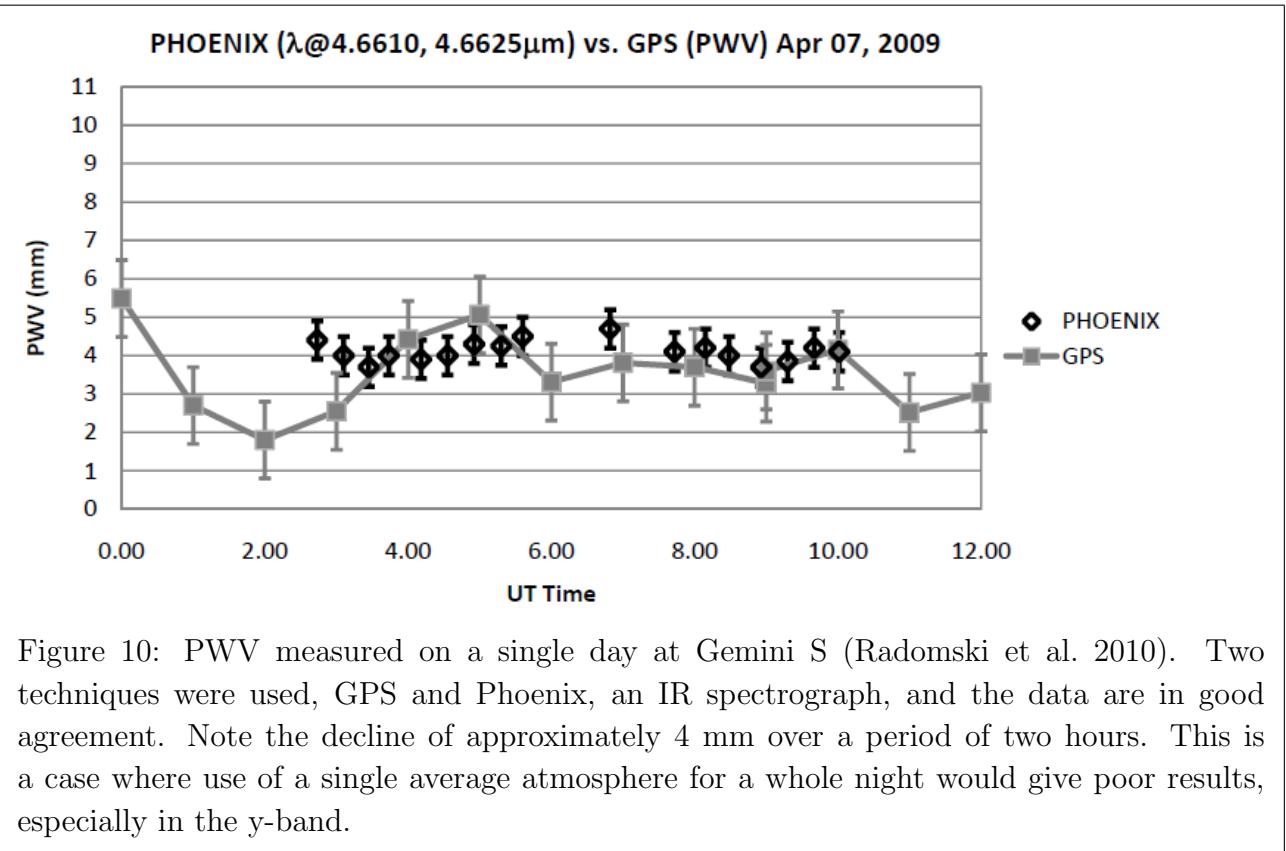


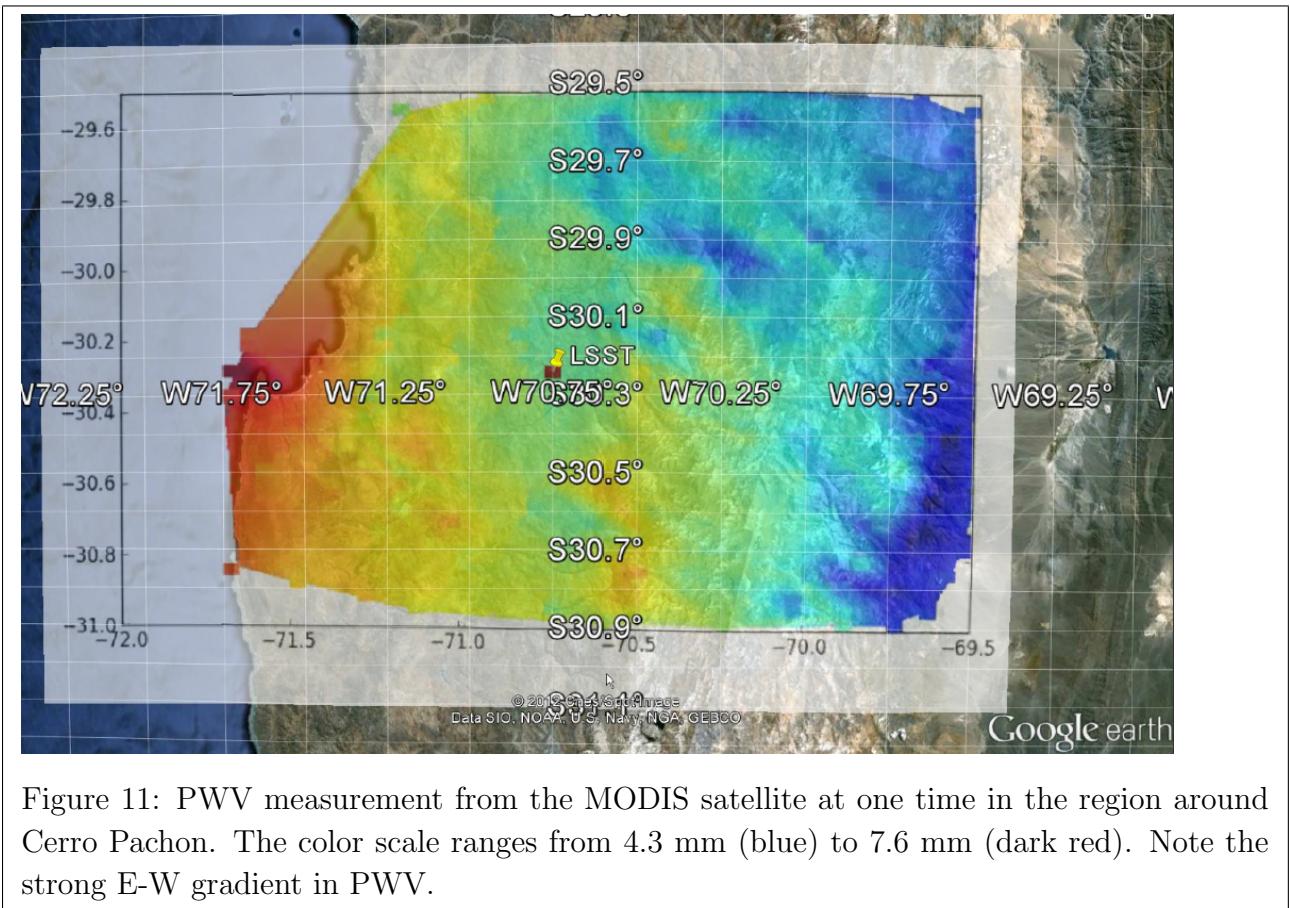


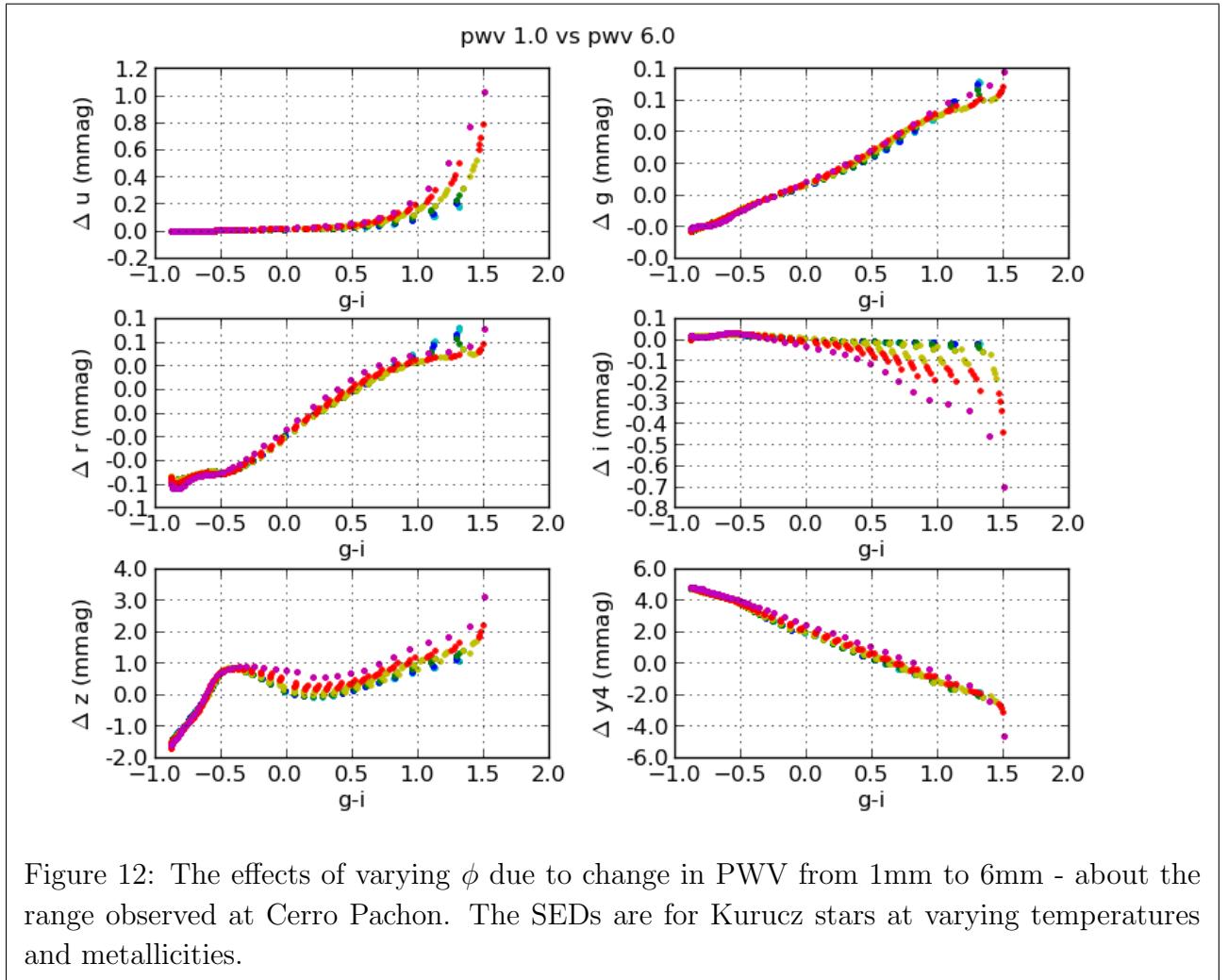












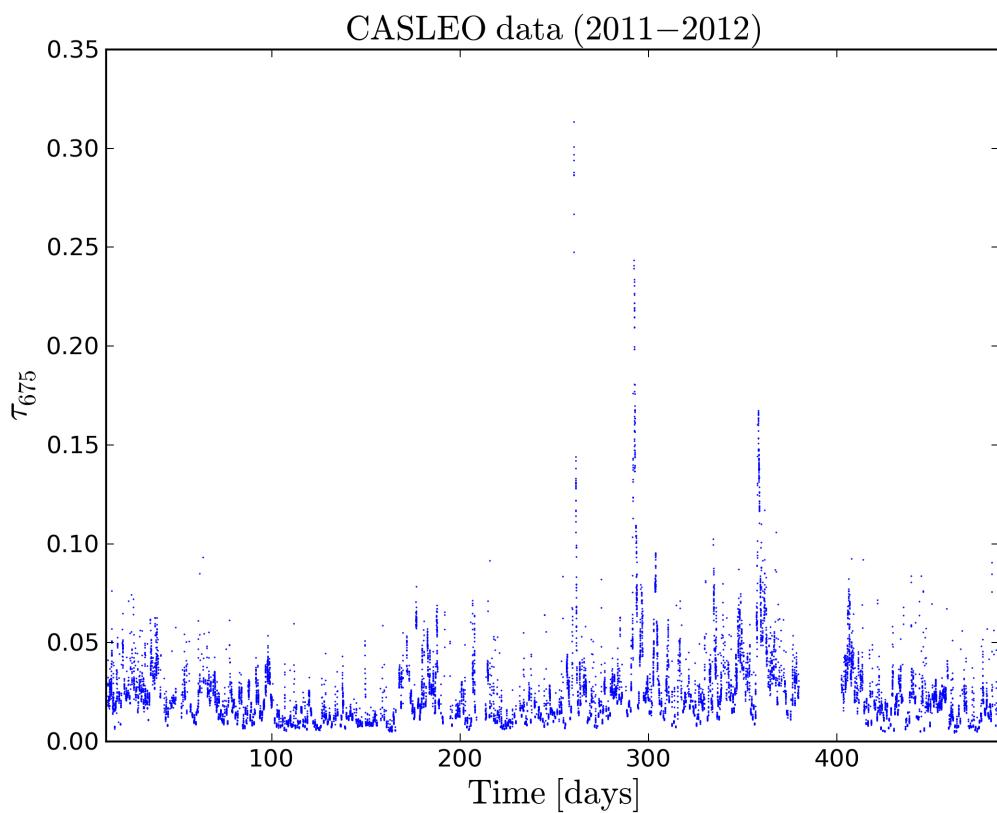
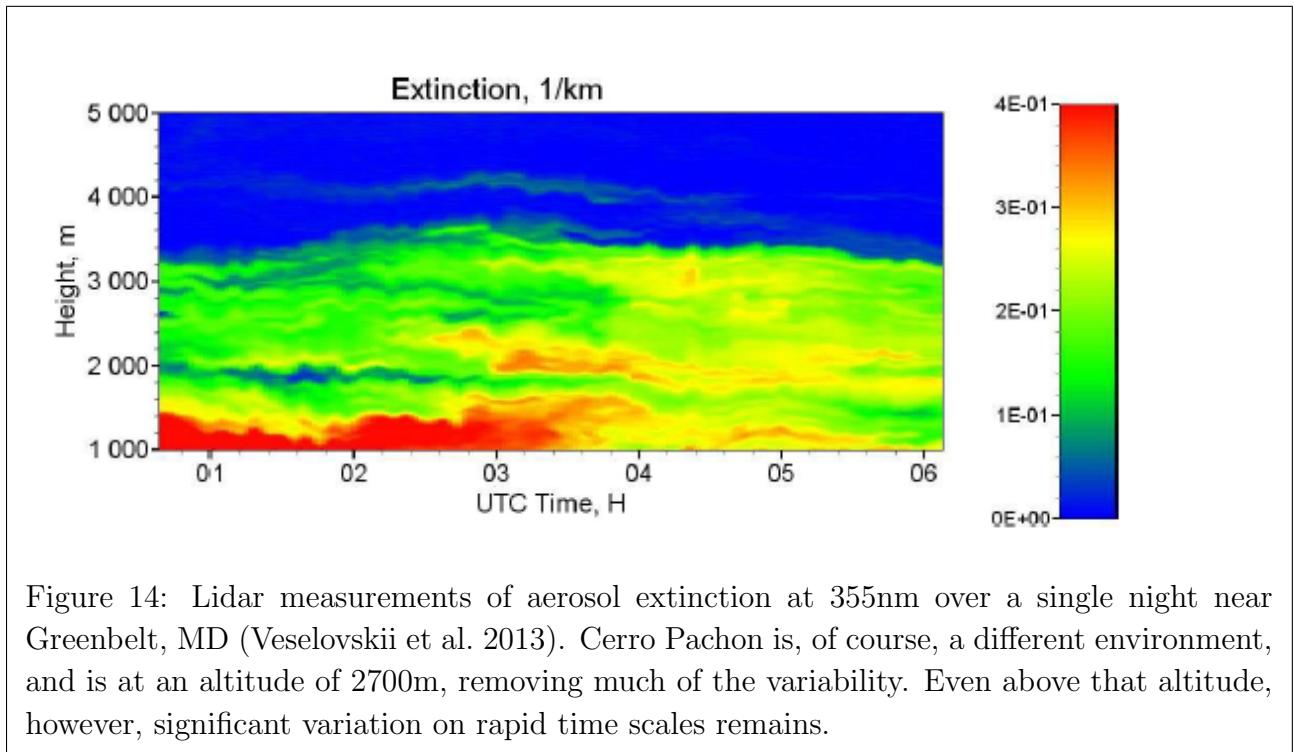
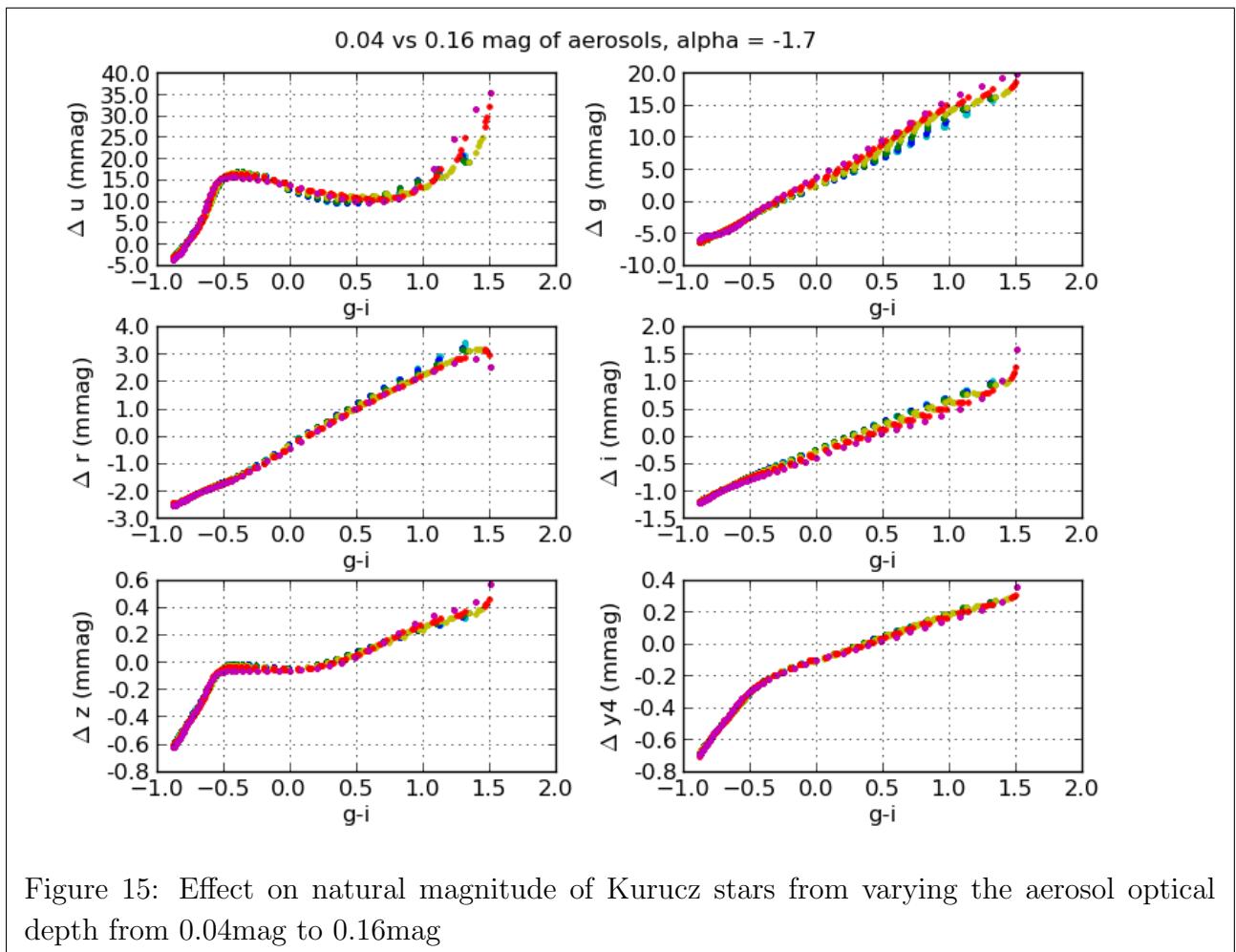
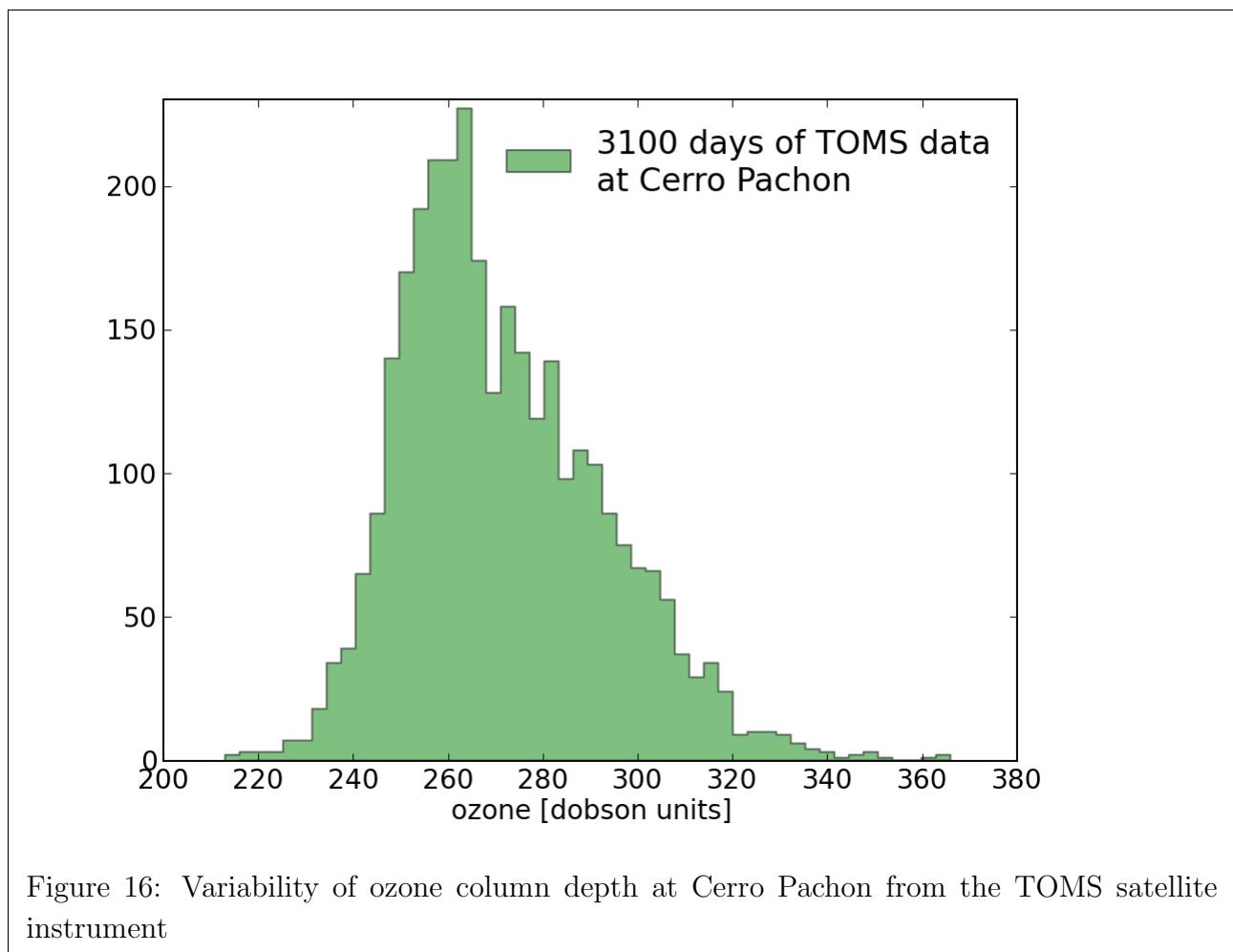
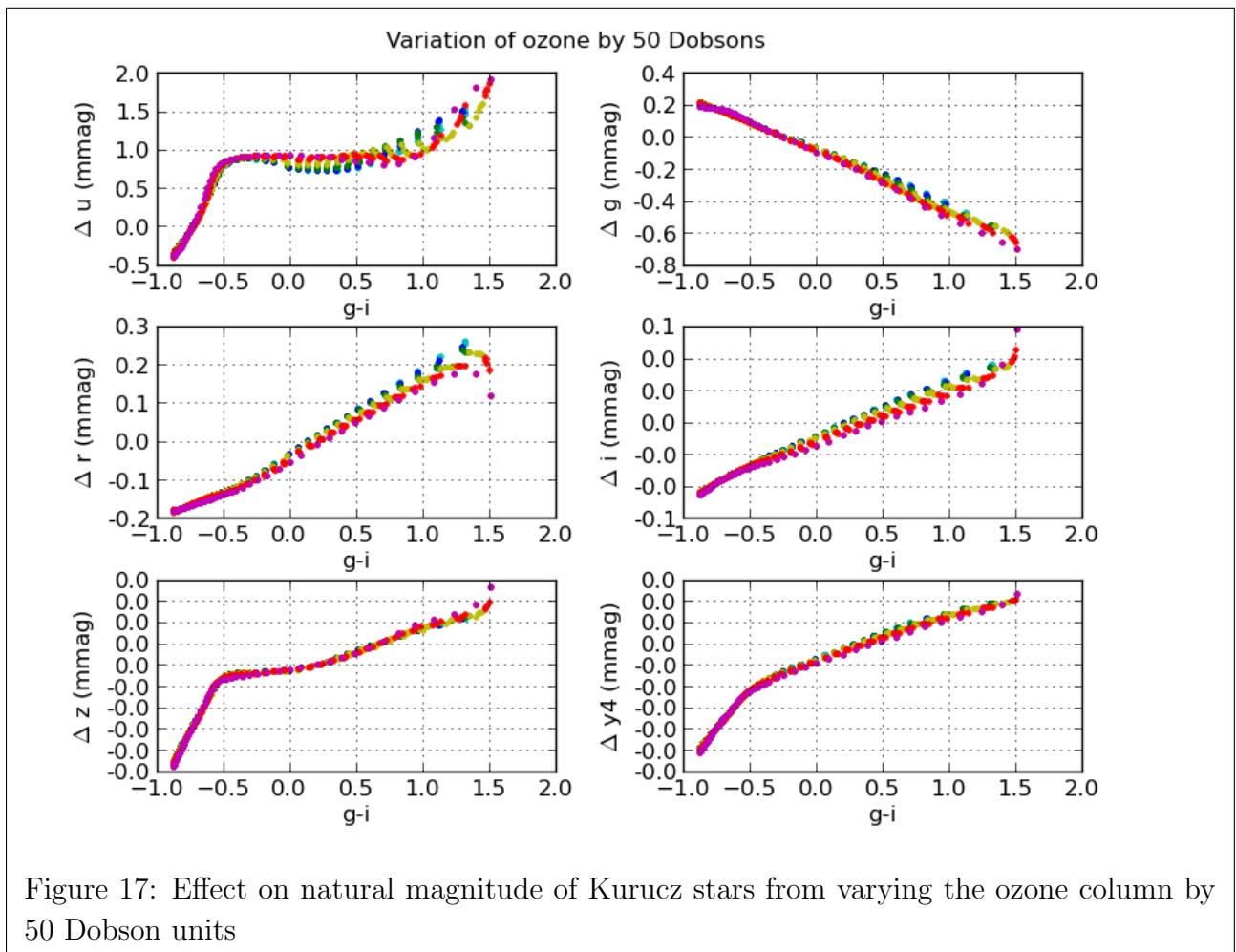


Figure 13: Varying aerosol optical depth at CASLEO, El Leoncito, Argentina. The site elevation at 2550m is similar to CP.









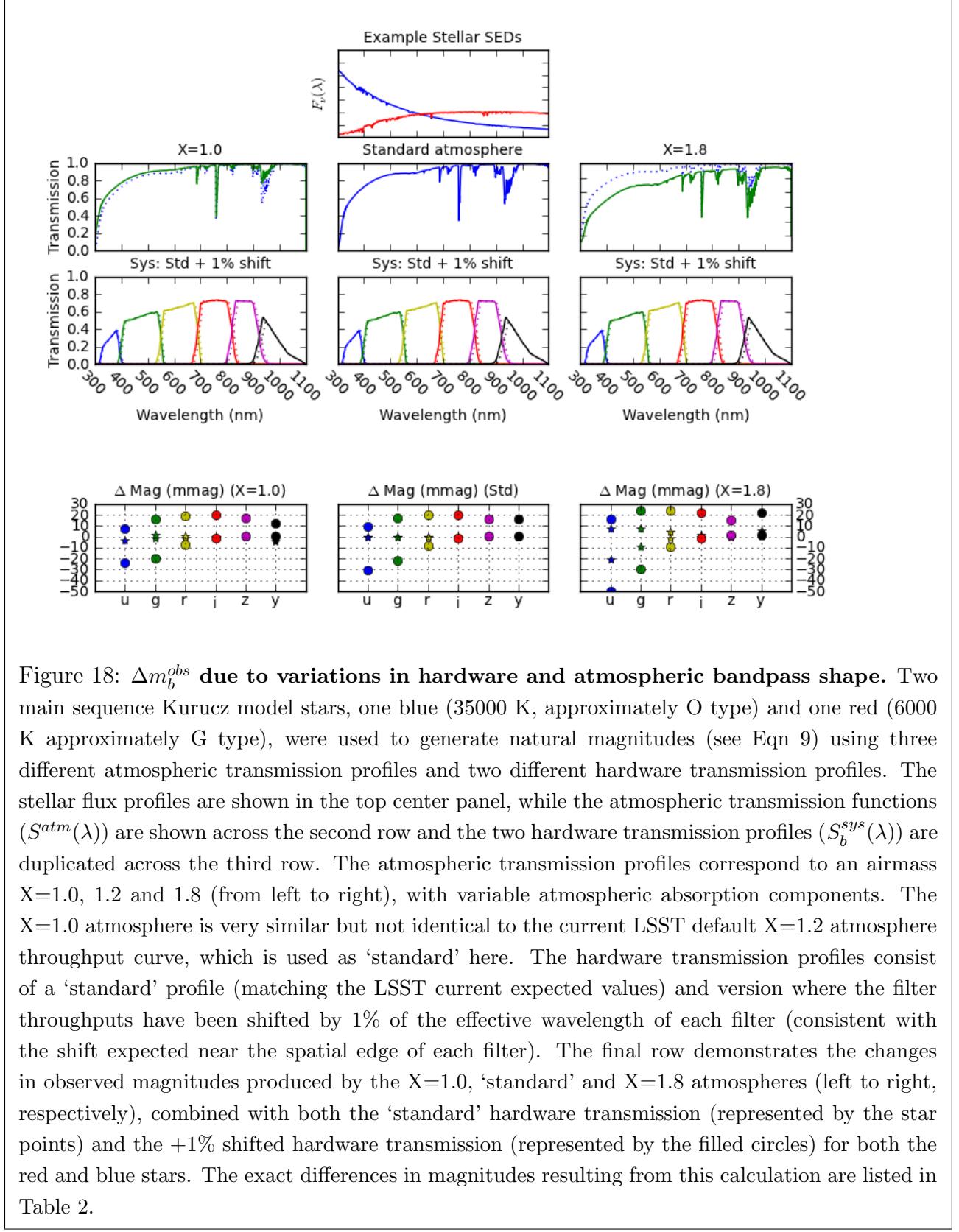
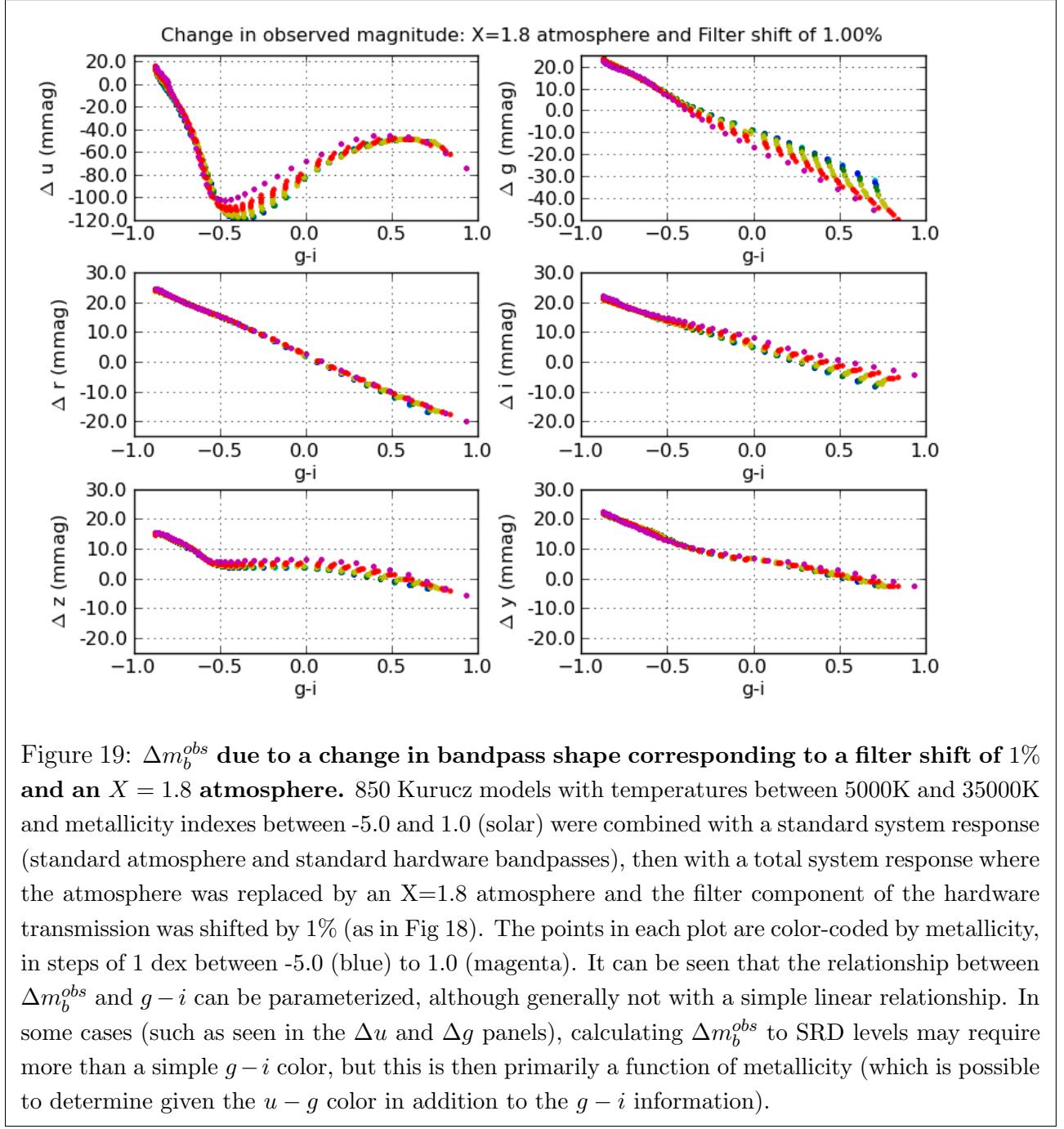
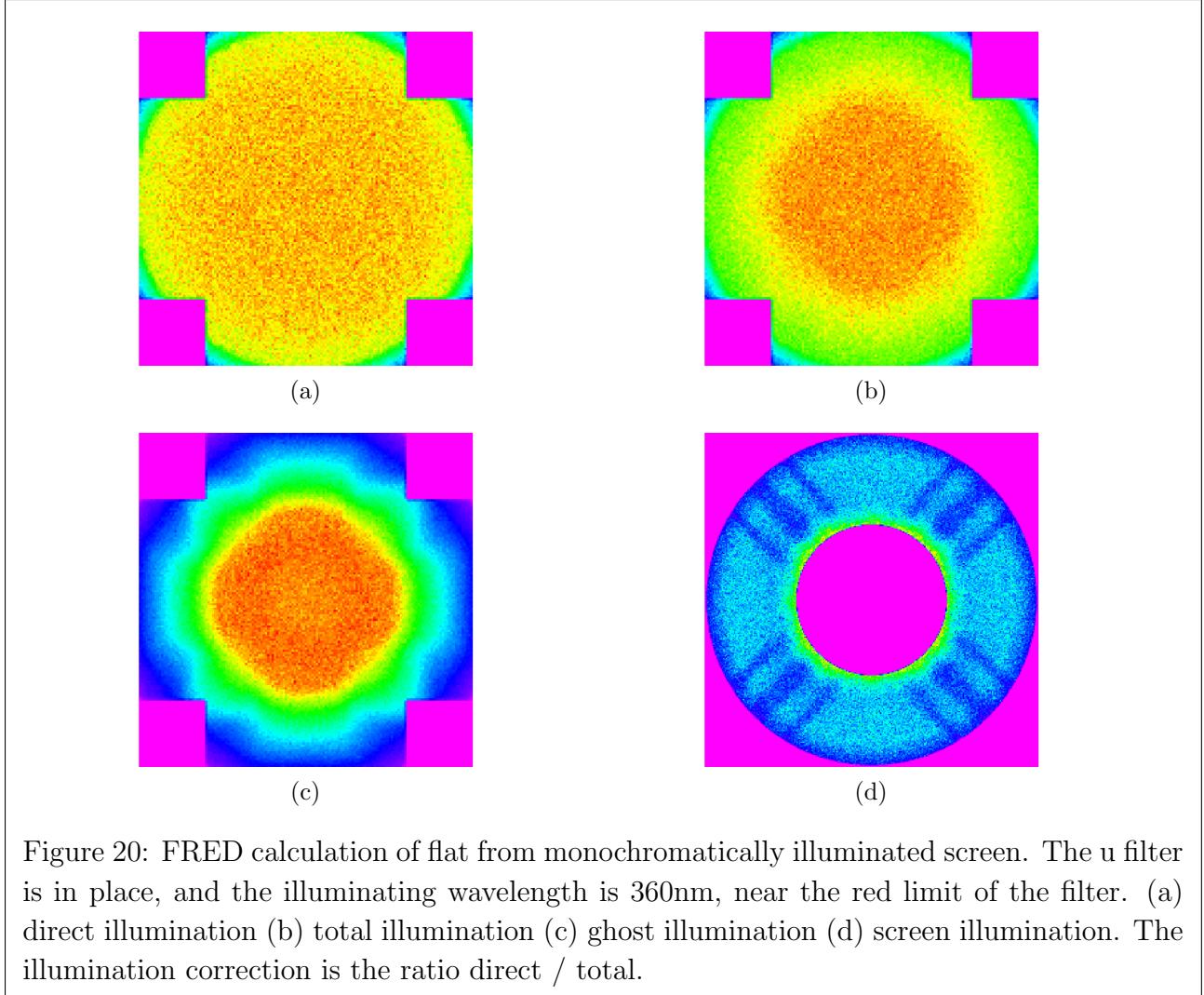


Table 2::  $\Delta m_b^{obs}$  due to variations in system and atmospheric bandpass shape (see also Fig 18). The first two rows show the baseline ('standard') magnitude of the star. All other rows show the *change* in magnitude (in mmag) due to the variations listed at left. Any value larger than 5 mmag would be larger than the RMS scatter allowed by the SRD. *TODO color-code values larger than 5 mmag*

| Bandpass                   | star | <i>u</i> (mag)    | <i>g</i>   | <i>r</i>   | <i>i</i>   | <i>z</i>   | <i>y</i>   |
|----------------------------|------|-------------------|------------|------------|------------|------------|------------|
| Std (X=1.2) atm, std sys   | red  | 21.472            | 20.378     | 20.000     | 19.911     | 19.913     | 19.913     |
| Std (X=1.2) atm, std sys   | blue | 19.102            | 19.503     | 20.000     | 20.378     | 20.672     | 20.886     |
|                            |      | $\Delta u$ (mmag) | $\Delta g$ | $\Delta r$ | $\Delta i$ | $\Delta z$ | $\Delta y$ |
| Std (X=1.2), +1% sys shift | red  | -31               | -22        | -8         | -2         | 1          | 1          |
| Std (X=1.2), +1% sys shift | blue | 9                 | 17         | 20         | 20         | 16         | 16         |
| X=1.0, std sys             | red  | 7                 | 2          | 0          | 0          | -0         | -1         |
| X=1.0, std sys             | blue | -3                | -1         | -1         | -0         | 1          | -4         |
| X=1.0, +1% sys shift       | red  | -24               | -20        | -8         | -1         | 1          | 0          |
| X=1.0, +1% sys shift       | blue | 7                 | 16         | 19         | 20         | 18         | 12         |
| X=1.8, std sys             | red  | -21               | -10        | -2         | -0         | 0          | 1          |
| X=1.8, std sys             | blue | 8                 | 8          | 4          | 2          | -1         | 6          |
| X=1.8, +1% sys shift       | red  | -50               | -30        | -10        | -2         | 1          | 2          |
| X=1.8, +1% sys shift       | blue | 16                | 24         | 24         | 22         | 15         | 22         |





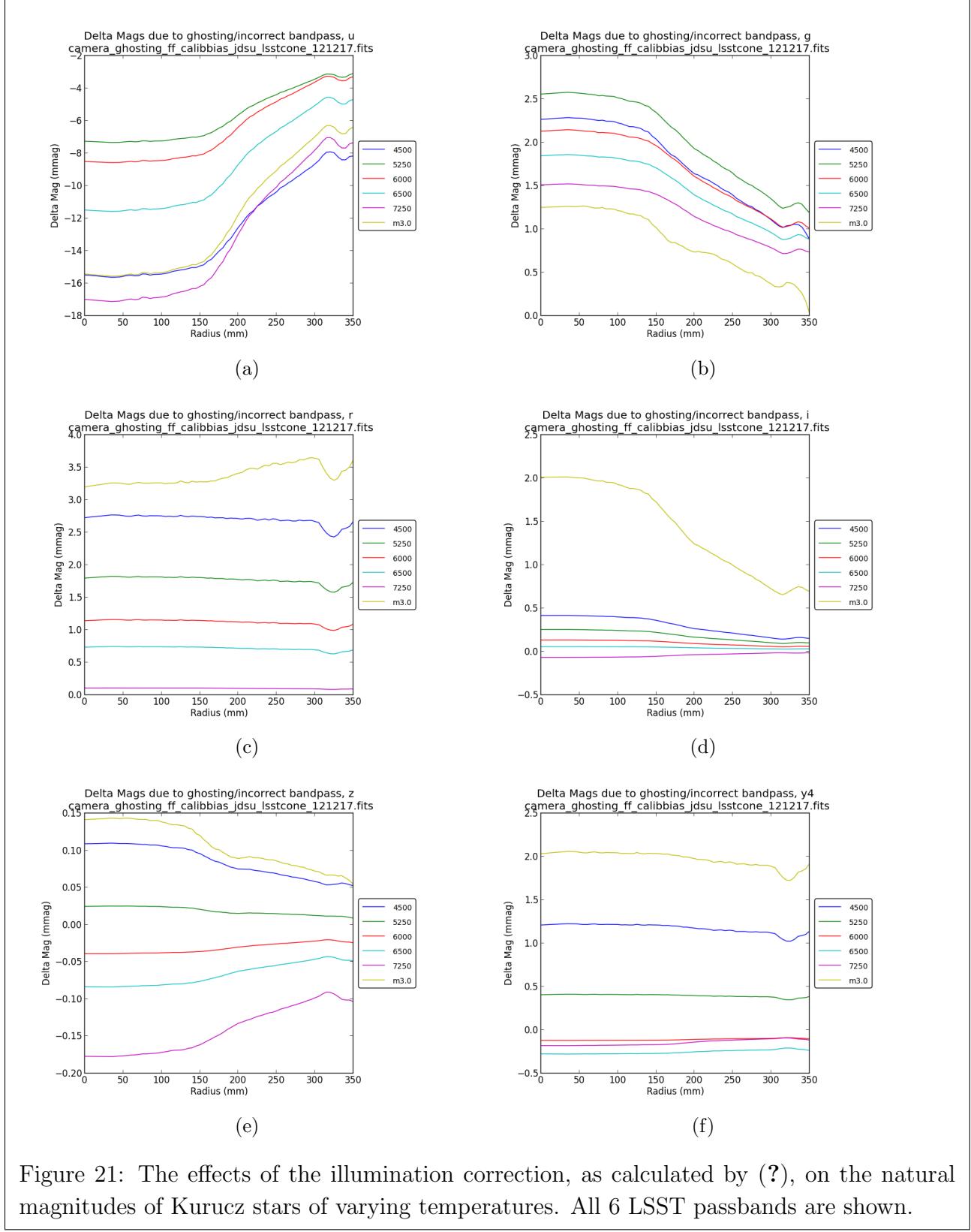


Figure 21: The effects of the illumination correction, as calculated by (?), on the natural magnitudes of Kurucz stars of varying temperatures. All 6 LSST passbands are shown.

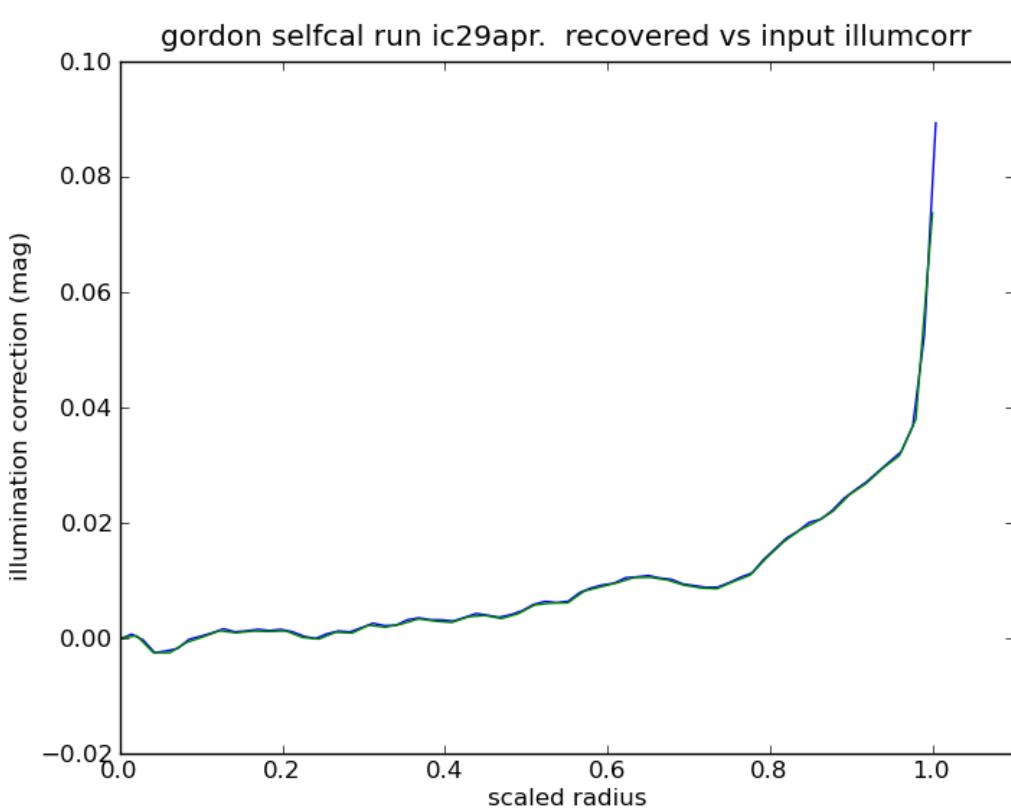


Figure 22: Recovery of input broadband illumination correction by self calibration. The input illumination correction was wavelength independent, but strongly dependent on radial position, as shown. The recovered illumination correction is essentially identical to that input.

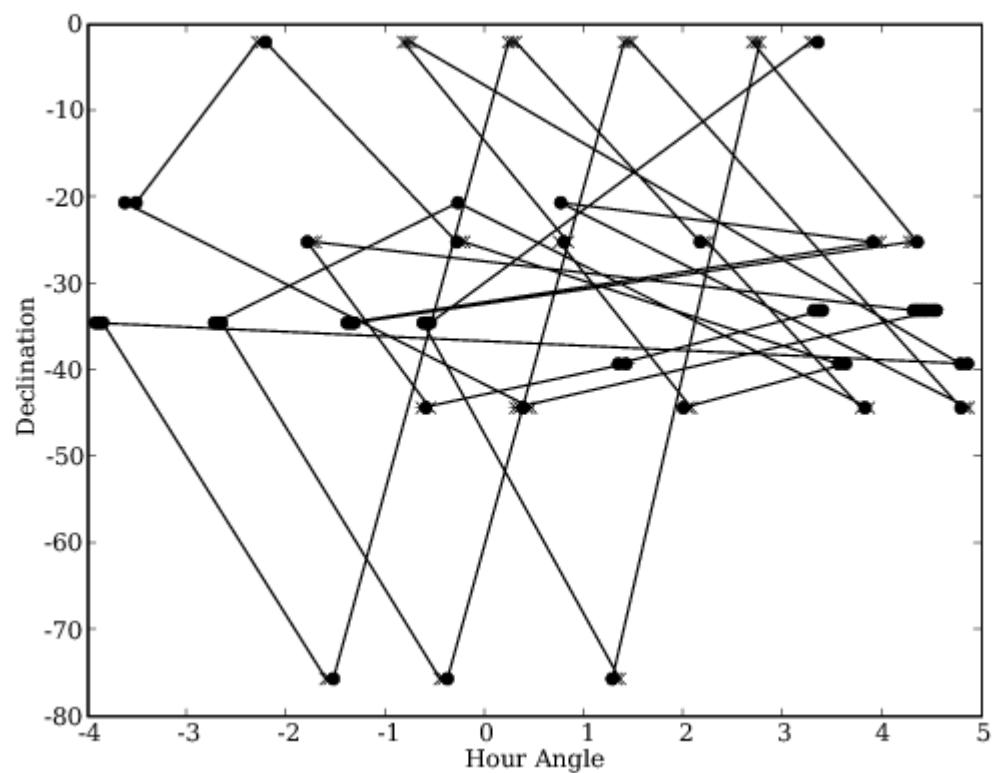
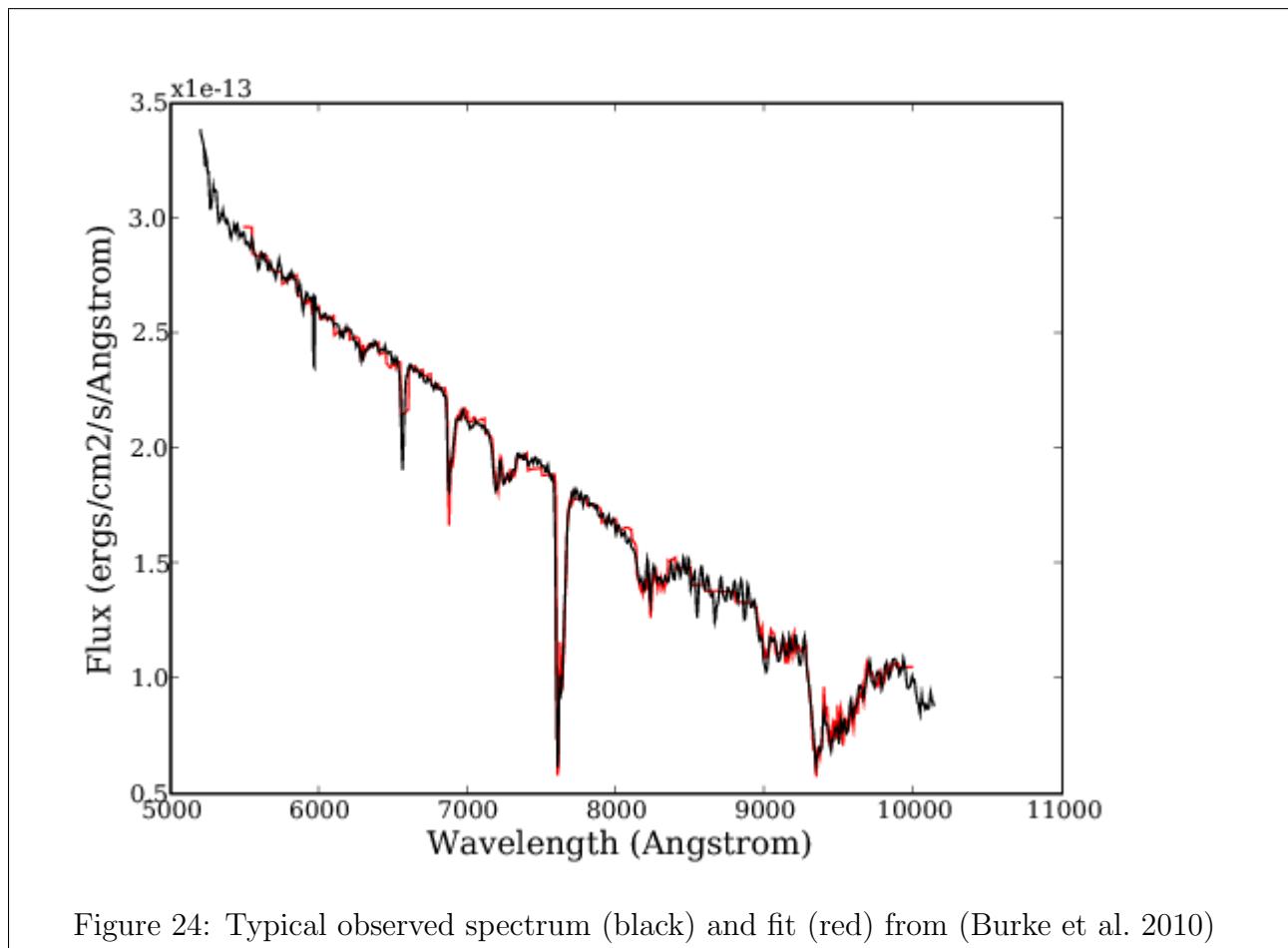


Figure 23: Pattern of observing atmospheric probe stars from (Burke et al. 2010). The Solid lines trace the temporal order of the observations



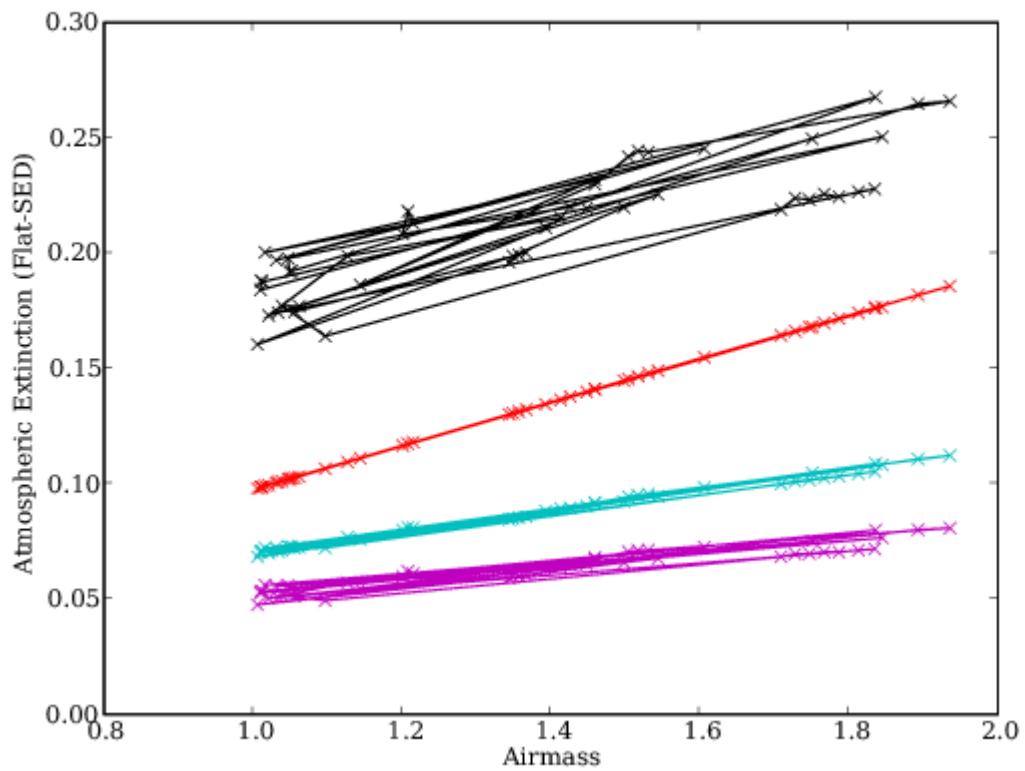


Figure 25: Non-gray atmospheric extinction from (Burke et al. 2010). Red is r-band, cyan is i-band, magenta is z-band, black is y-band.

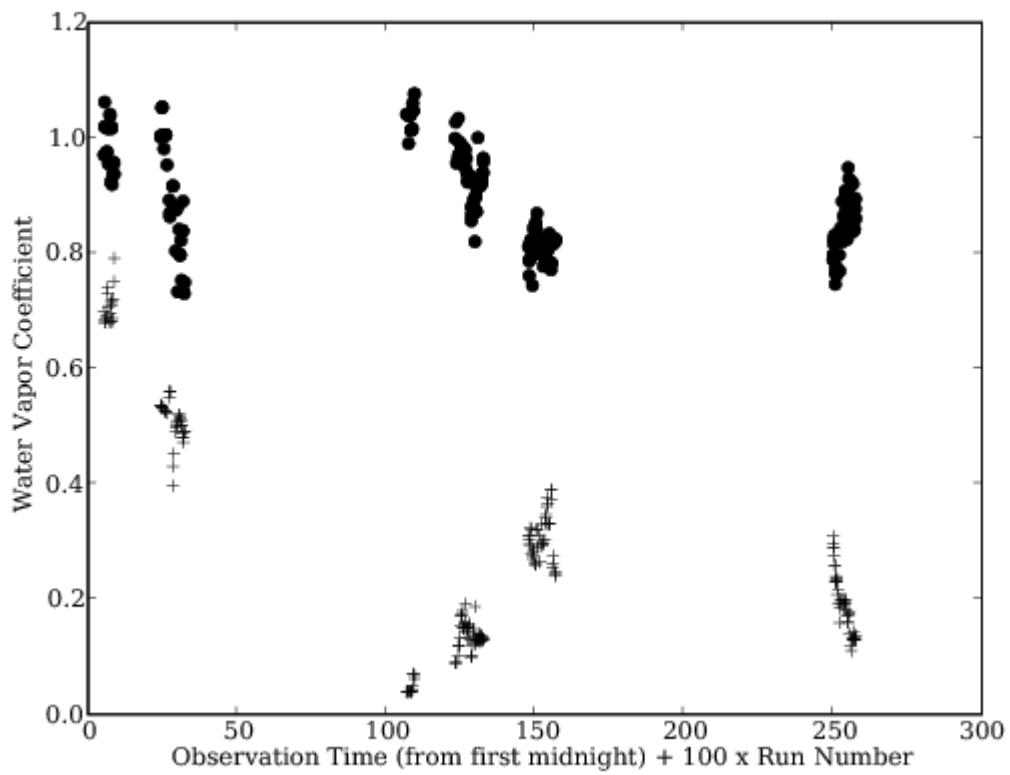
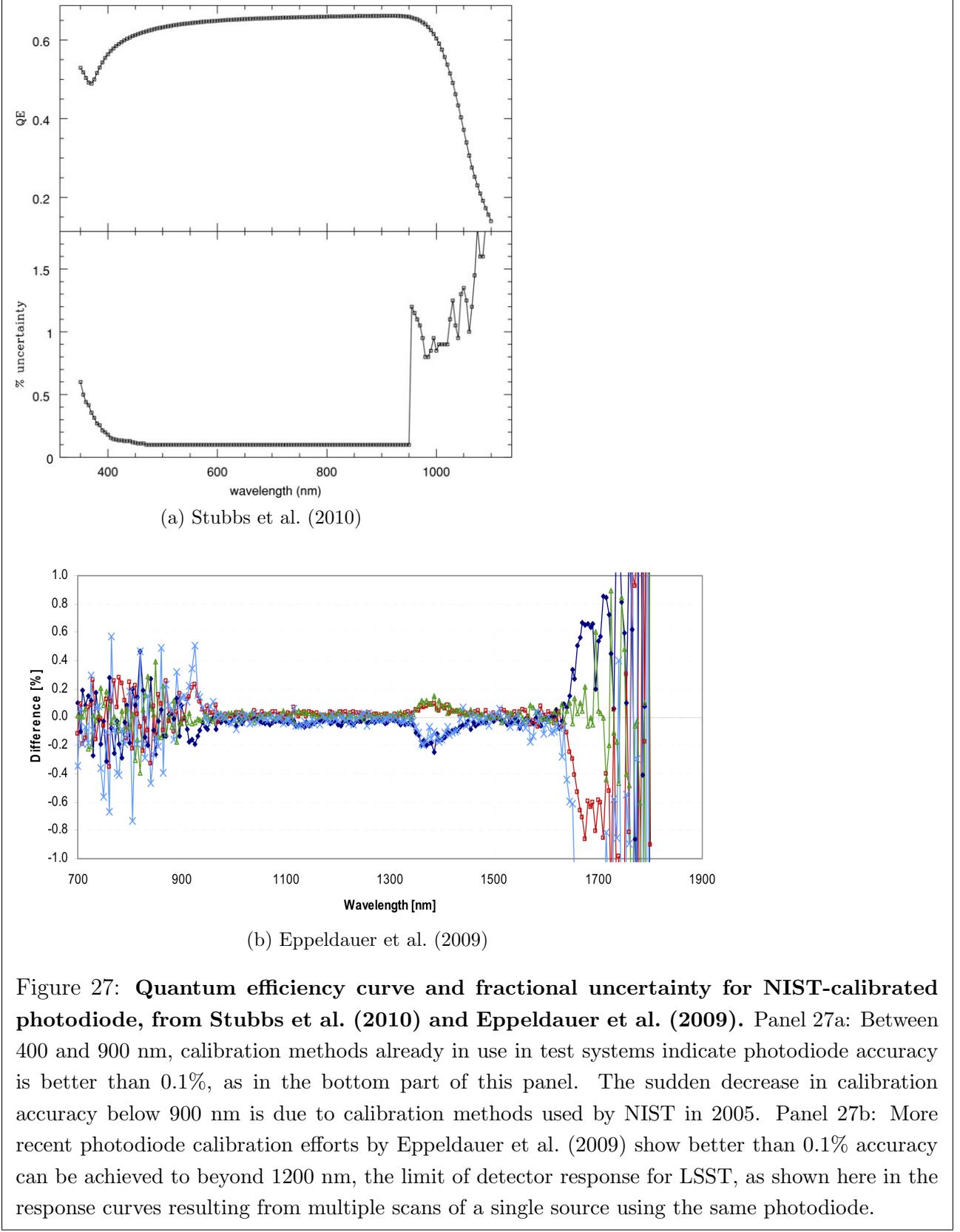


Figure 26: Fitted water vapor coefficient from (Burke et al. 2010), expressed as a ratio to the standard value (filled circles). Crosses are relative humidity from CTIO (changed scale).



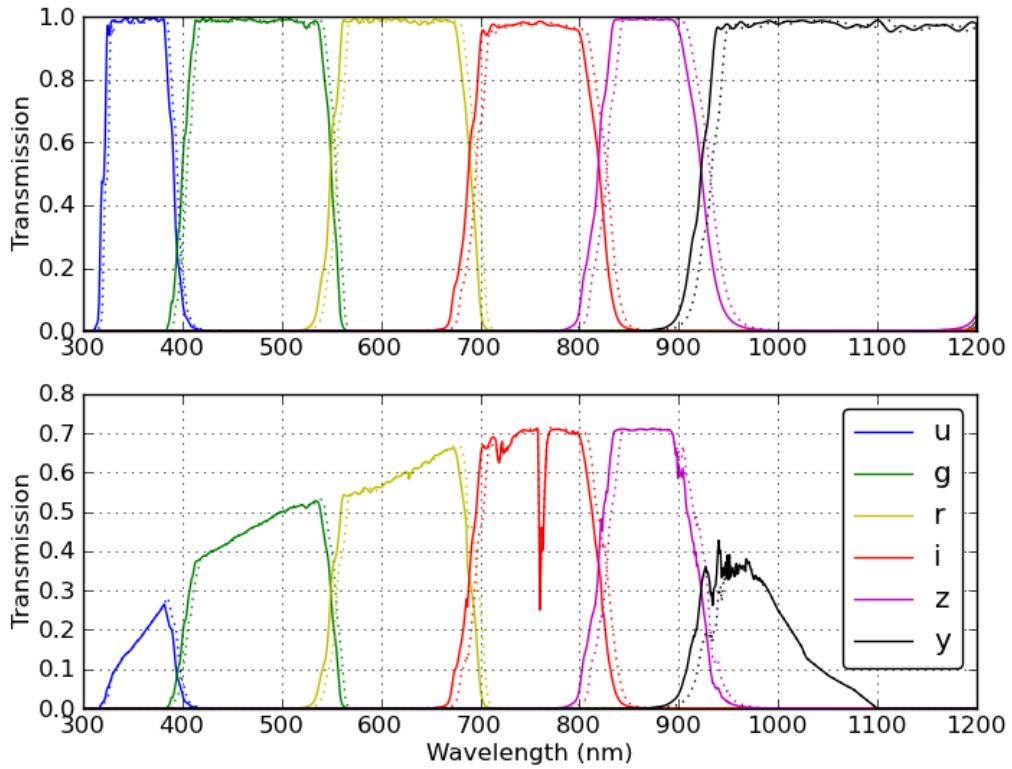
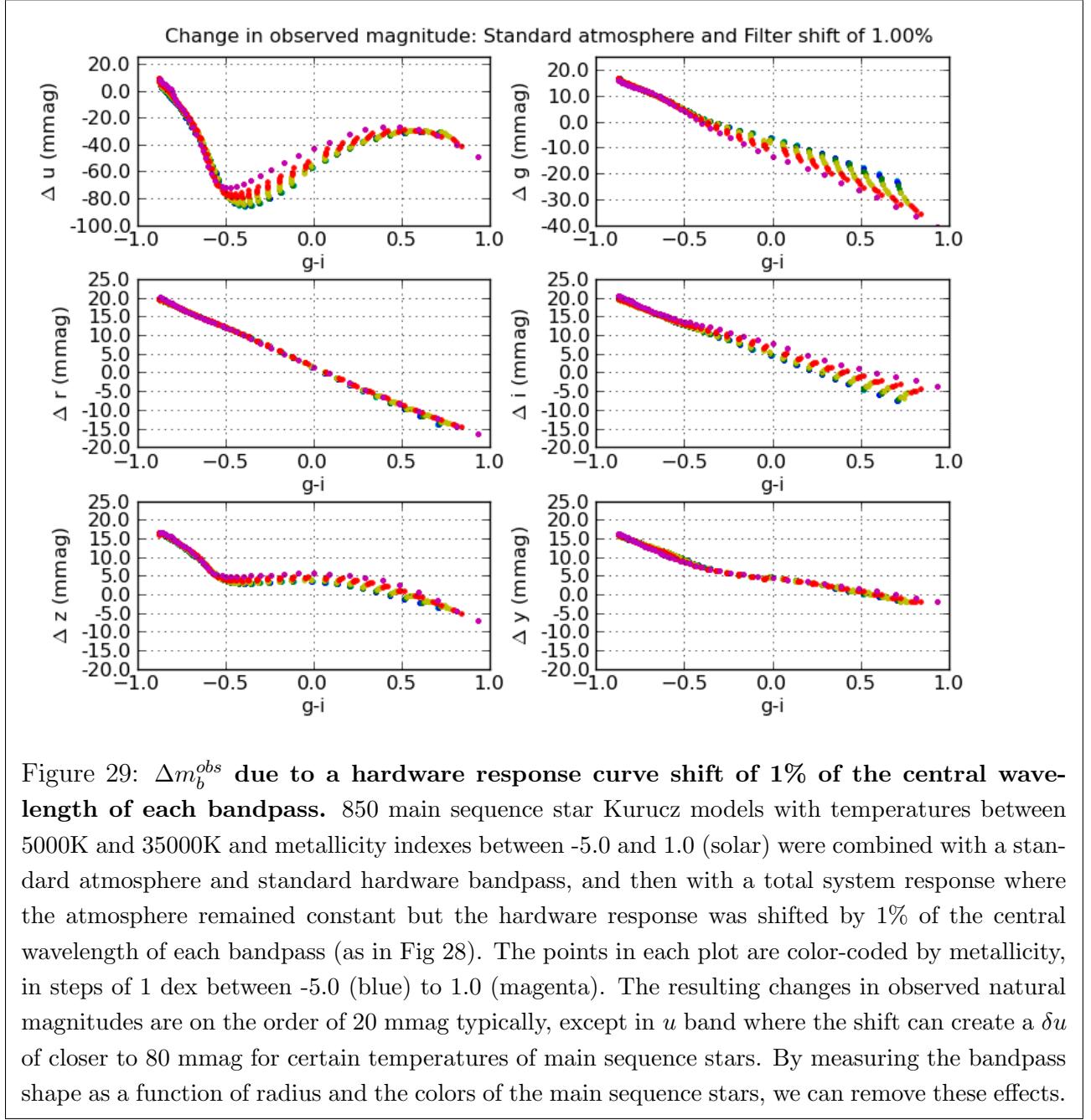
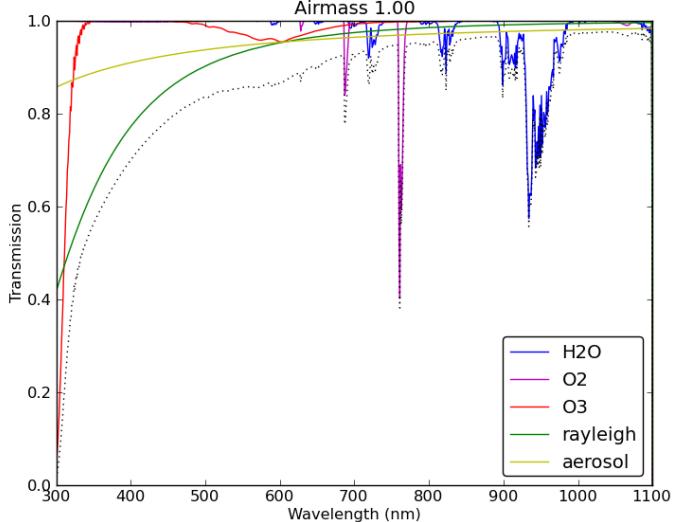
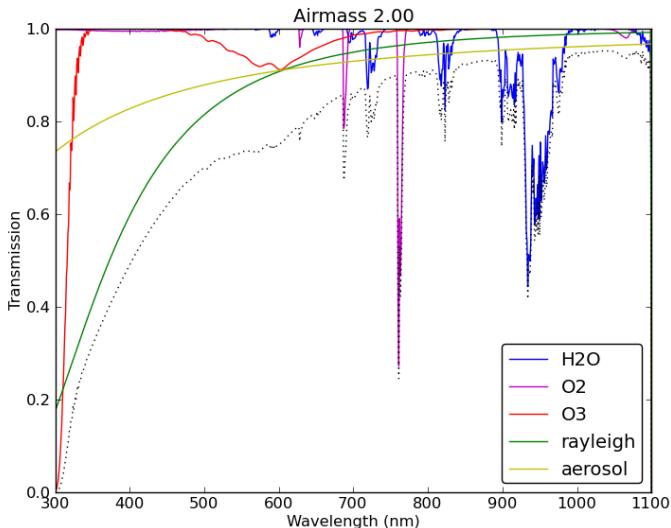


Figure 28: **Baseline filter curves and a potential (1% of the central wavelength) shift due to nonuniformity in the filter bandpass.** The solid lines indicate standard filter bandpasses (top panel: filter alone, bottom panel: filter plus standard mirror, lens, detector and atmosphere response curves) while the dashed lines indicate the same bandpass shifted redward by 1% of the central wavelength.



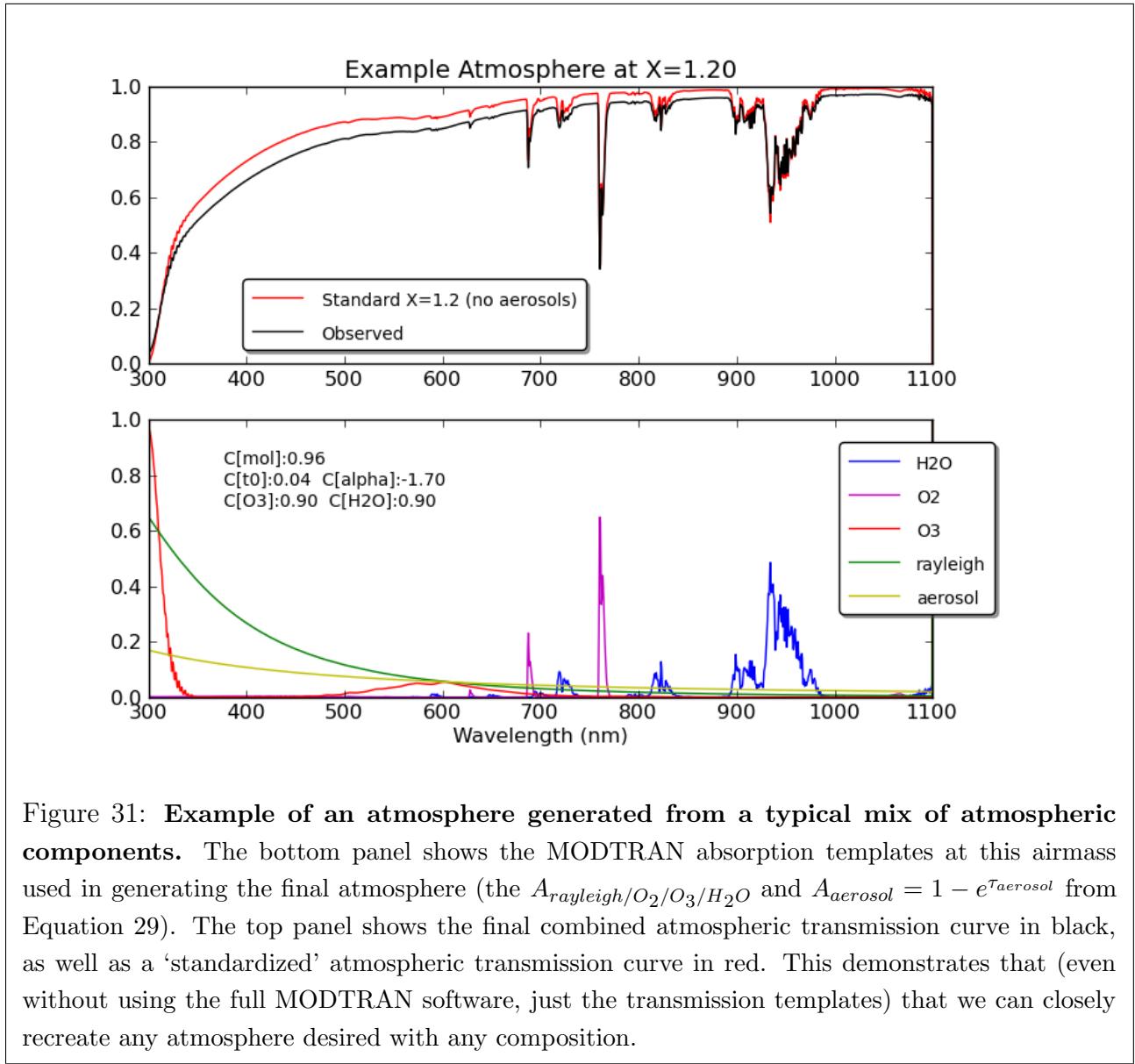


(a)



(b)

**Figure 30: Components of atmospheric absorption.** The wavelength dependence of various atmospheric absorption components at zenith (Panel 30a) and at airmass=2.0 (Panel 30b) are shown here. The H<sub>2</sub>O (blue) and O<sub>3</sub> (red) molecular absorption contributions are shown separately, while the O<sub>2</sub> absorption is combined with other trace elements (magenta). A typical example of aerosol scattering (Mie scattering) is included (yellow), as is molecular scattering (Rayleigh scattering) (green). All components except aerosol scattering were generated using MODTRAN4 with the US Standard option (aerosol scattering is not part of the US Standard atmosphere). The resulting total absorption curve is the product of each of these effects and is shown with the dotted black line. This is an illustrative atmosphere; under actual observing conditions the molecular absorption components will vary in strength with time and the square root of the airmass, the molecular and aerosol scattering will depend on airmass, and the aerosol scattering profile will also vary with time.



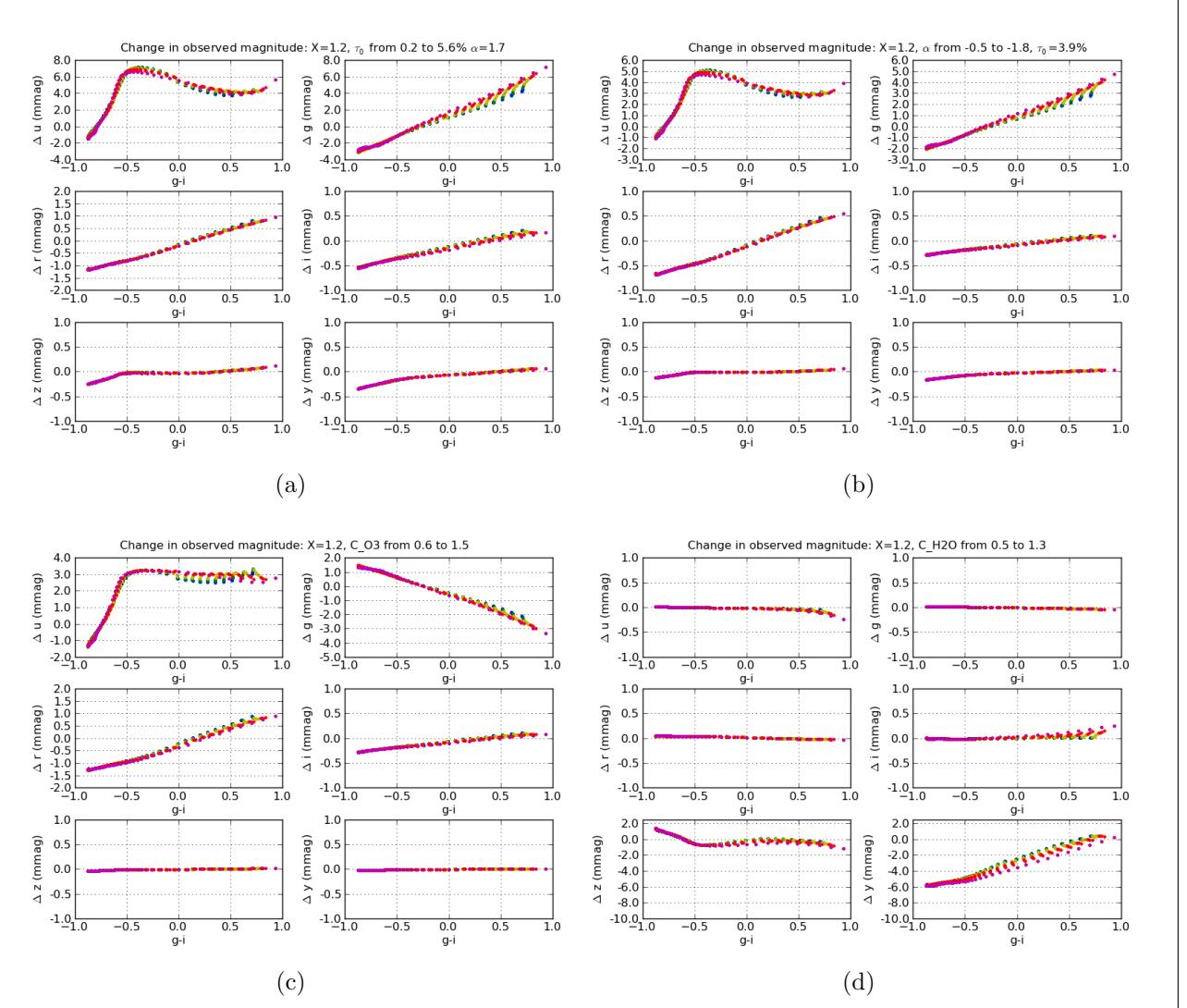
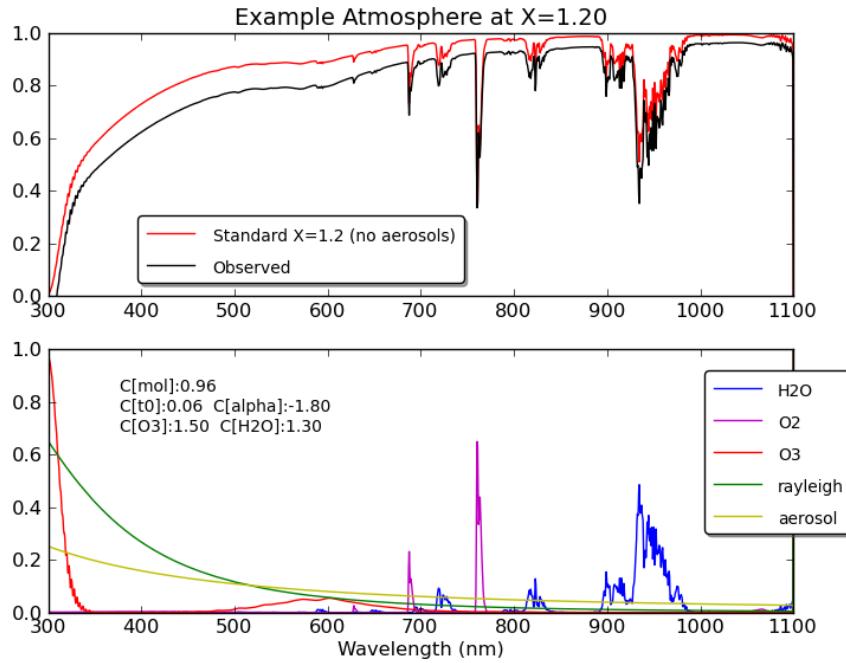
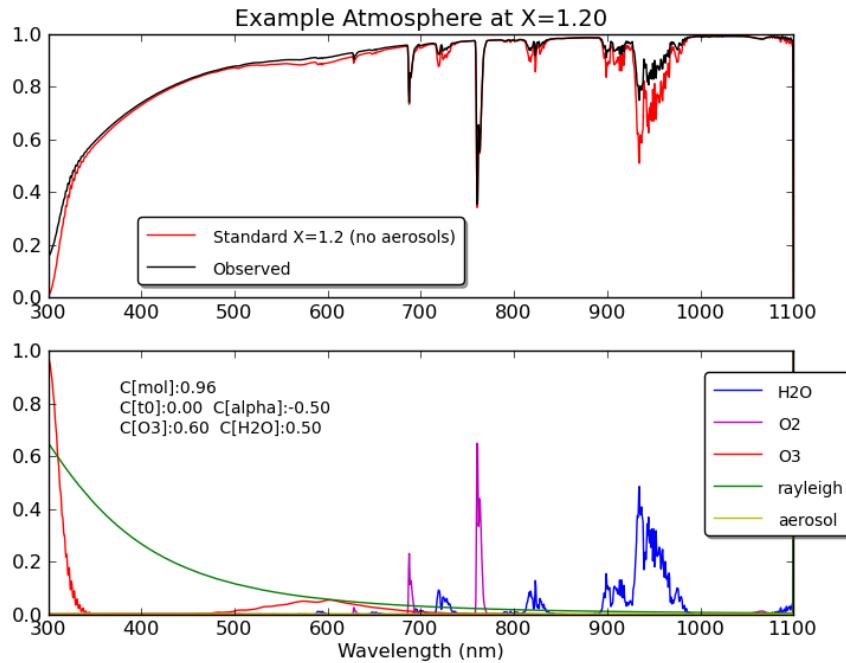


Figure 32:  $\Delta m_b^{obs}$  due to variations of each individual absorption component. Each atmospheric transmission curve (at  $X=1.2$ ) was combined with the set of main sequence Kurucz curves to determine the resulting changes in observed magnitudes, as in Figure 29. Panels 39a and 39b show the effects of varying aerosol absorption in  $\tau_0$  and  $\alpha$  respectively, Panel 32c shows the effect of varying  $O_3$  absorption. These effects are concentrated in  $u$  and  $g$  bands, with a negligible effect in  $izy$ . Panel 32d shows the effect of varying the  $H_2O$  absorption, which is strongest in  $y$ , with some effect in  $z$  and no effect in  $ugri$ .

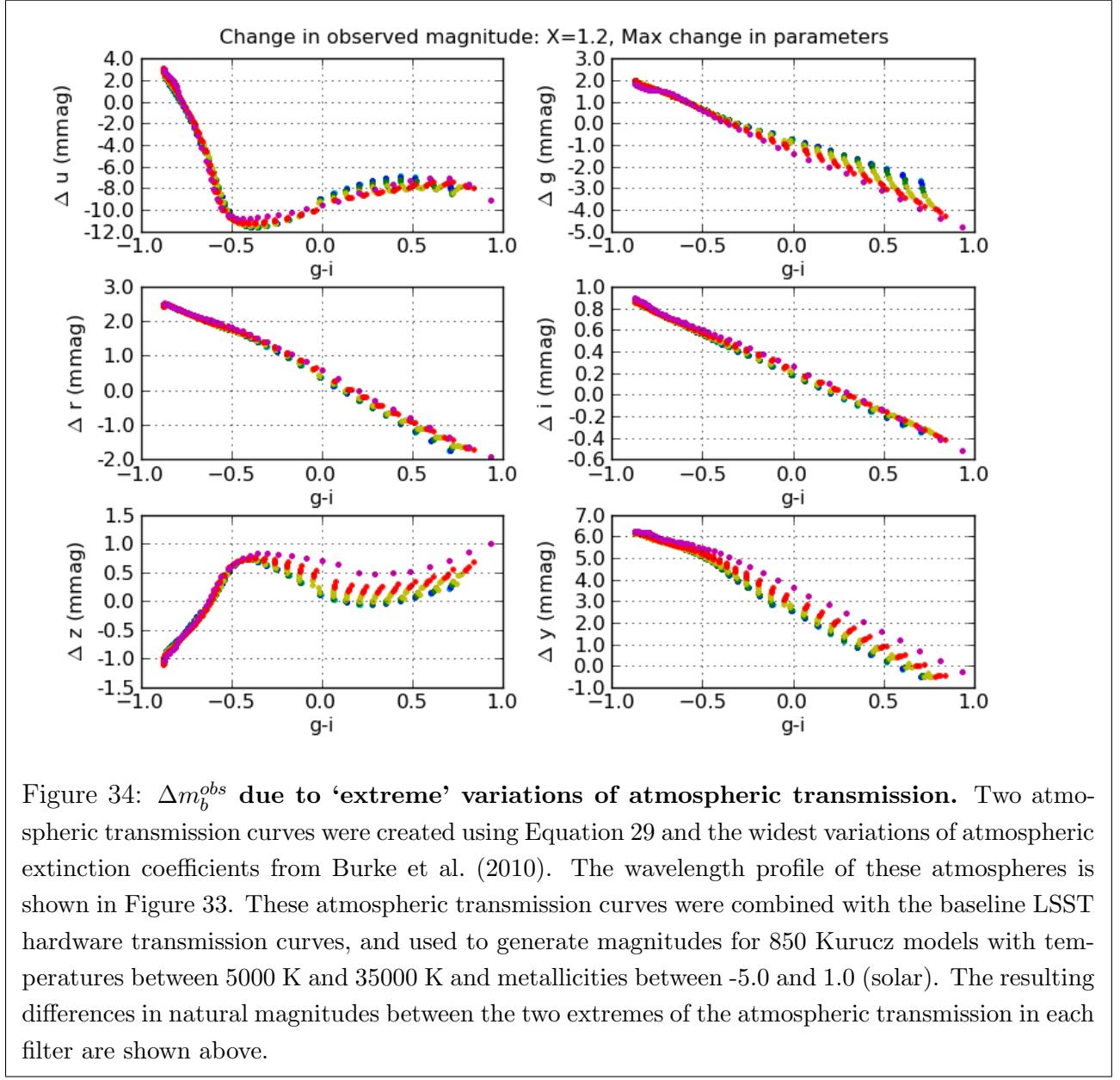


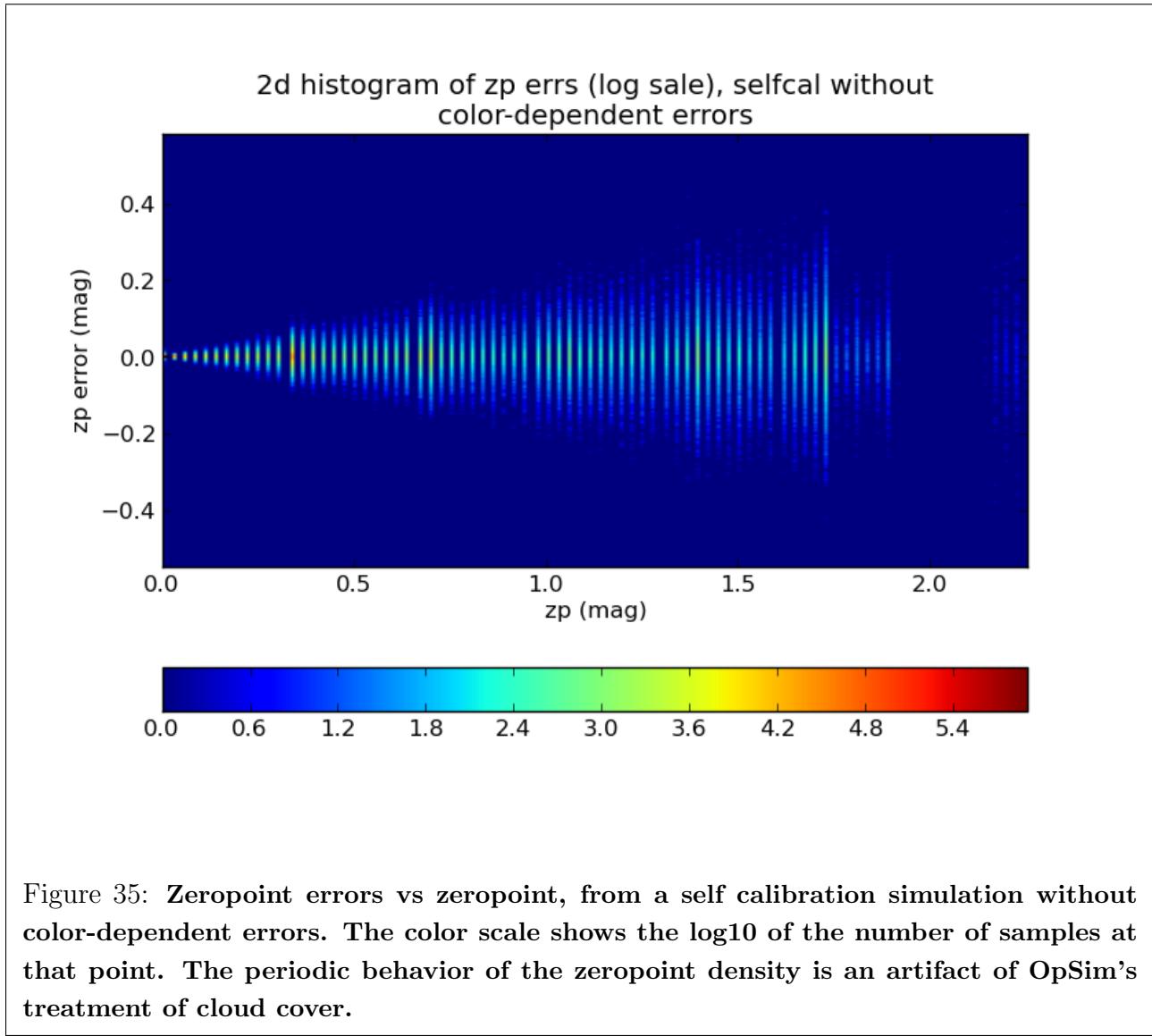
(a)

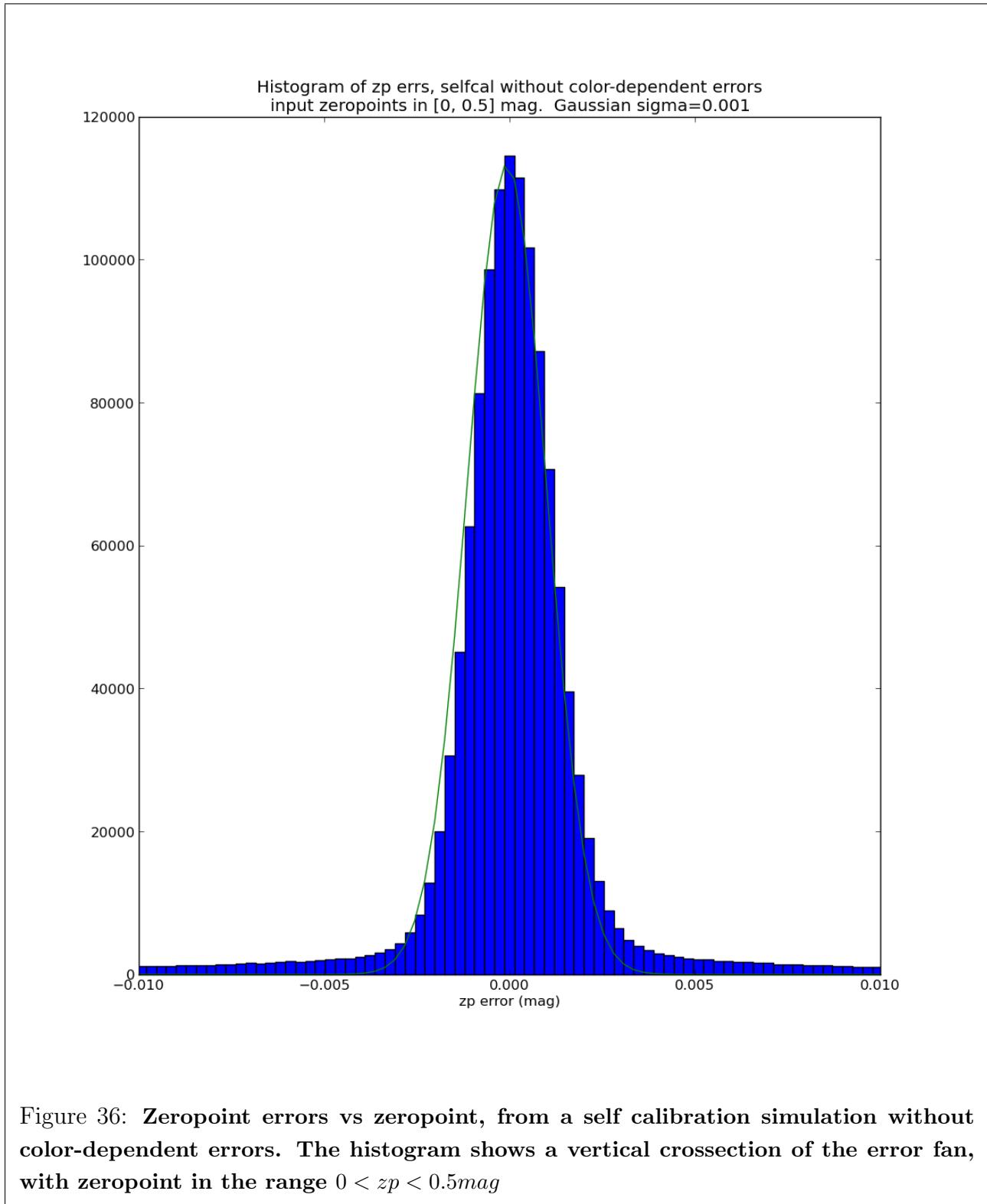


(b)

Figure 33: ‘Extreme’ atmospheres generated from MODTRAN profiles and extremes of atmospheric coefficients. Using the extremes of  $C_{H_2O}$ ,  $C_{O_3}$ , and  $\tau_0$  and  $\alpha$  from Burke et al. (2010), two test atmospheres with  $X = 1.2$  were created using Equation 29.







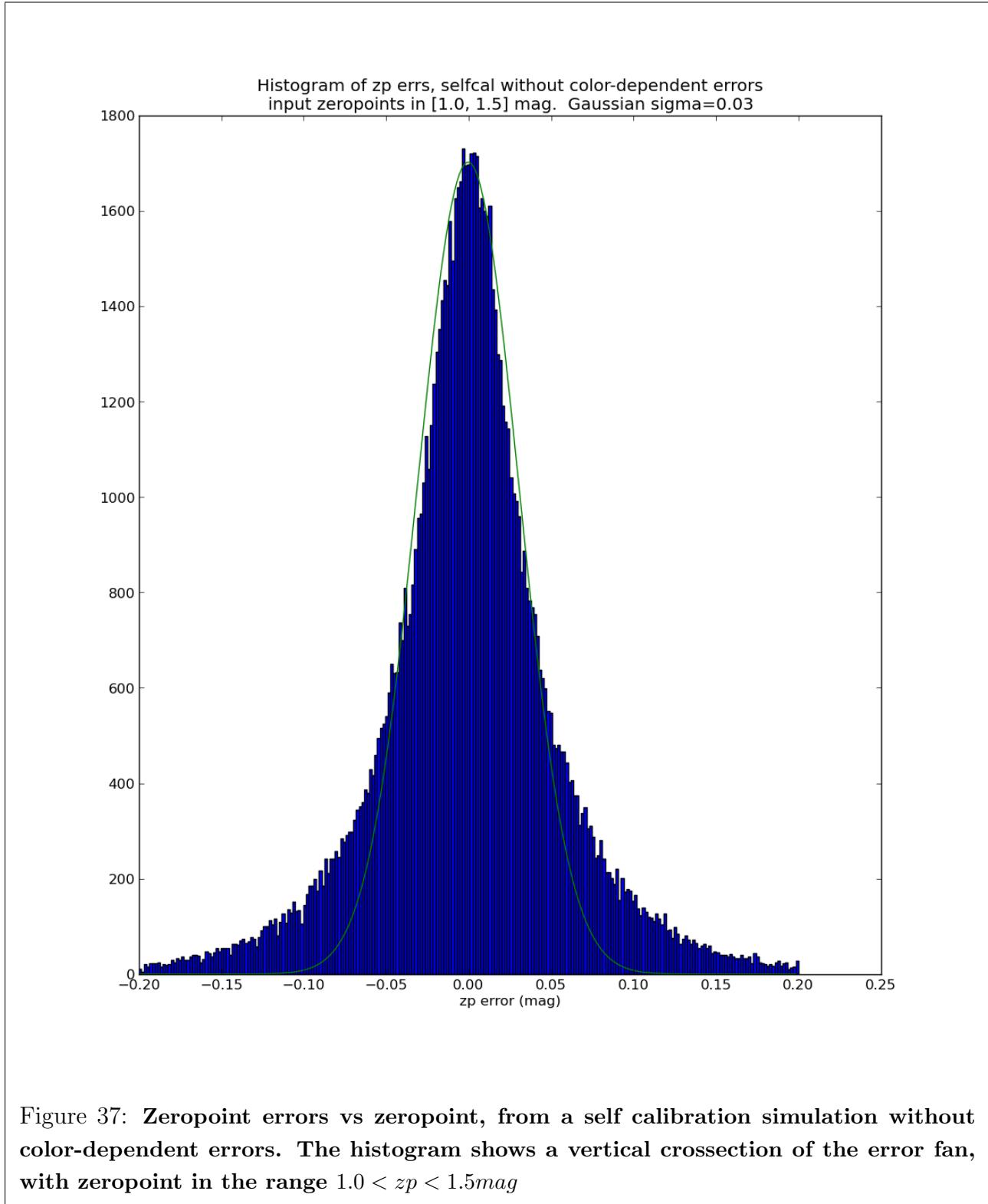




Table 3:: Repeatability error budget. All values are in mmag

| Affected term             | Effect                                       | <i>u</i> | <i>g</i> | <i>r</i> | <i>i</i> | <i>z</i> | <i>y</i> |
|---------------------------|--|----------|----------|----------|----------|----------|----------|
| $m_b^{obs}$               | <b>Total</b>                                 | 3.0      | 3.0      | 3.0      | 3.0      | 3.0      | 3.0      |
| $\Delta m_b^{obs}$        | Atmospheric water vapor errors               | 0.2      | 0        | 0        | 0        | 1.0      | 2.0      |
|                           | Atmospheric aerosol and ozone errors         | 1.6      | 2.1      | 0.1      | 0        | 0        | 0        |
|                           | Undetected atmospheric variability           | 1.4      | 0.7      | 0        | 0        | 0        | 0        |
|                           | Monochromatic illumination correction errors | 3.0      | 0.5      | 0.8      | 0.5      | 0.1      | 0.6      |
|                           | Photodiode monitoring system errors          | 0.8      | 0.5      | 0        | 0        | 0        | 0.1      |
|                           | Calibration star SED errors                  | 5.6      | 0.8      | 0.5      | 0.4      | 0.4      | 0.2      |
|                           | Focal plane temperature errors               | 0        | 0        | 0        | 0        | 0        | 0.2      |
|                           | <b>Total</b>                                 | 6.7      | 2.5      | 0.9      | 0.6      | 1.1      | 2.1      |
| $Z_b^{obs}$               | Clouds                                       |          |          |          |          |          |          |
|                           | Camera gain variation                        |          |          |          |          |          |          |
|                           | Shutter variation                            |          |          |          |          |          |          |
|                           | Temperature induced QE variations            |          |          |          |          |          |          |
|                           | Errors in standards                          |          |          |          |          |          |          |
|                           | <b>Total</b>                                 |          |          |          |          |          |          |
| <b>Total</b>              |  | 7.4      | 3.9      | 3.1      | 3.1      | 3.2      | 3.7      |
| <b>Design Requirement</b> |  | 7.5      | 5.0      | 5.0      | 5.0      | 7.5      | 7.5      |
| <b>Min. Requirement</b>   |  | 12       | 8.0      | 8.0      | 8.0      | 12       | 12       |

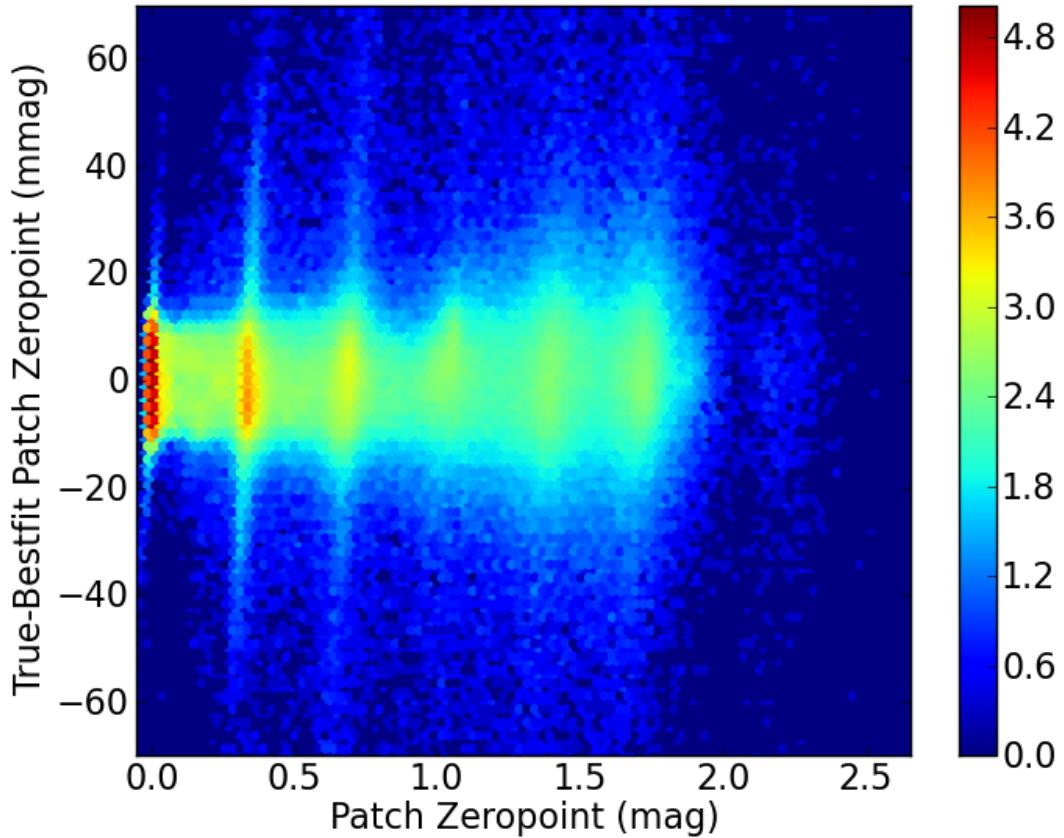


Figure 38: Zeropoint errors vs zeropoint, from a self calibration simulation. The color scale shows the  $\log_{10}$  of the number of samples at that point. The periodic behavior of the zeropoint density is an artifact of OpSim’s treatment of cloud cover.

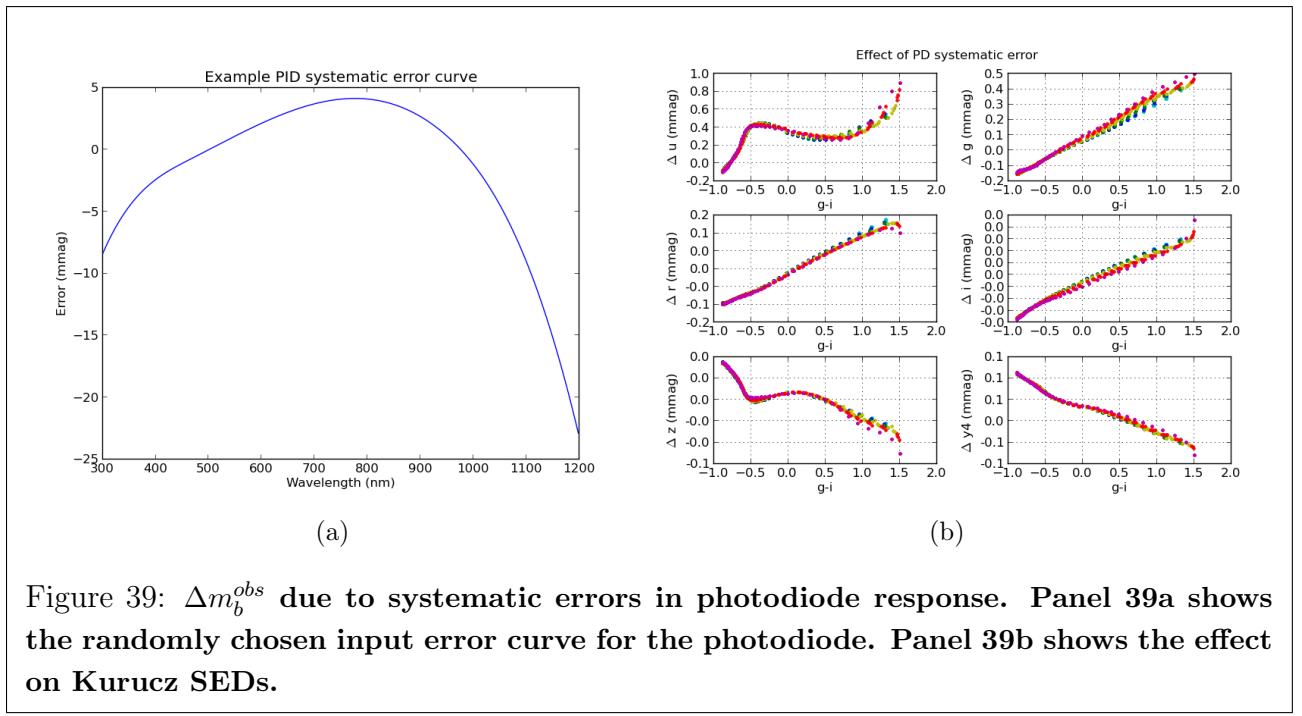


Figure 39:  $\Delta m_b^{obs}$  due to systematic errors in photodiode response. Panel 39a shows the randomly chosen input error curve for the photodiode. Panel 39b shows the effect on Kurucz SEDs.

