

Habitable Evaporated Cores

*Converting Mini-Neptunes into
Super-Earths in the Habitable
Zone of M Dwarfs*

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The Big Picture

M dwarfs: the best targets  Terrestrial planets are easiest to detect around low-mass stars

In situ formation unlikely  Rocky planets in the HZ could be small and volatile-poor
Raymond et al. (2007), Lissauer et al. (2007)

Planets can migrate  Gas-rich, volatile-rich planets can migrate into the HZ

Planets can lose mass  Roche lobe overflow and XUV-driven hydrodynamic escape
Erkaev et al. (2007), Lopez et al. (2012)

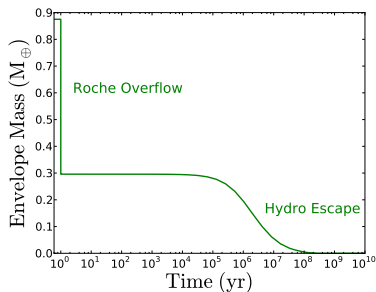
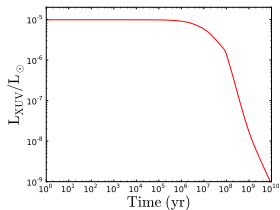
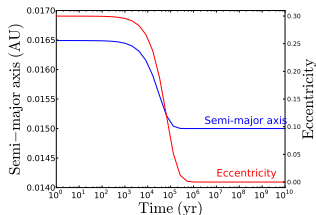
TESS  Observations of these planets are just around the corner



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The Model



Parameters

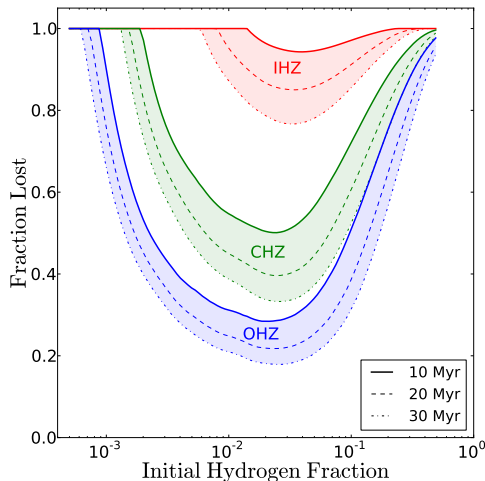
$M_p = 3.5M_{\oplus}$	$f_{H0} = 0.25$
$a_0 = 0.0165 \text{ AU}$	$e_0 = 0.3$
$M_* = 0.08M_{\odot}$	$\tau_{\text{CTL}} = 1 \text{ s}$
$L_{\text{sat}} = 10^{-3}L_*$	$\epsilon_{\text{XUV}} = 0.3$



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Forming a Habitable Planet



$$M_p = 3.5M_{\oplus}$$

$$e_0 = 0$$

$$M_* = 0.08M_{\odot}$$

$$\tau_{\text{CTL}} = 1 \text{ s}$$

$$L_{\text{sat}} = 10^{-3}L_*$$

$$\epsilon_{\text{XUV}} = 0.3$$

Gas-rich mini-Neptunes that migrate **early** into the **IHZ** form super-Earths.

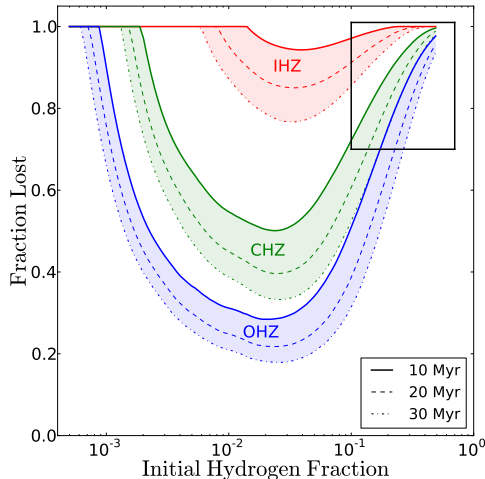
Let's take a closer look...



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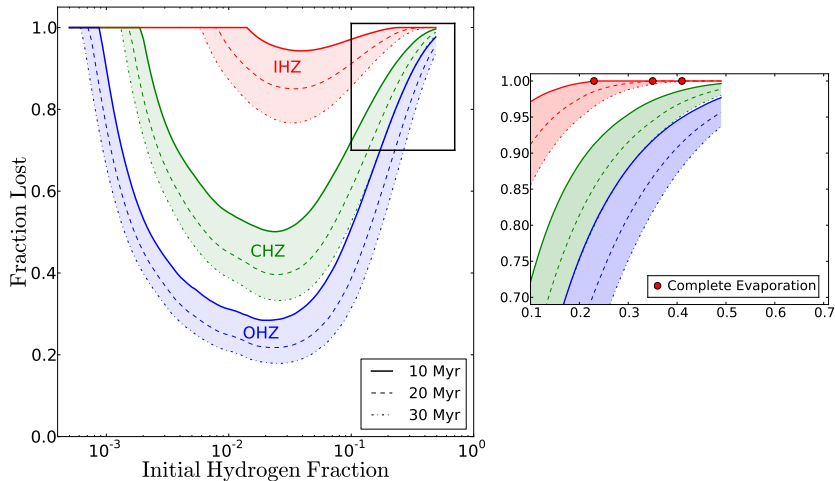


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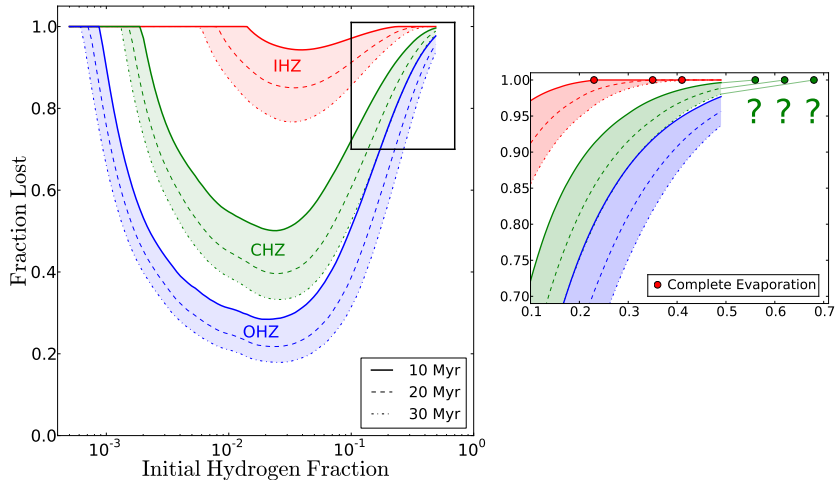
Forming a Habitable Planet



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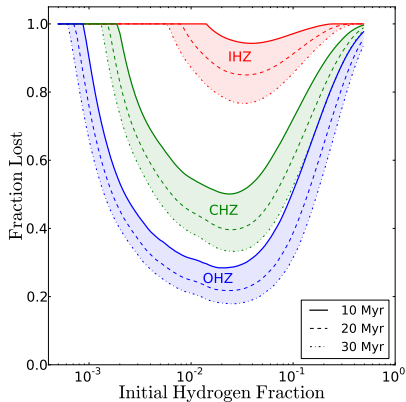
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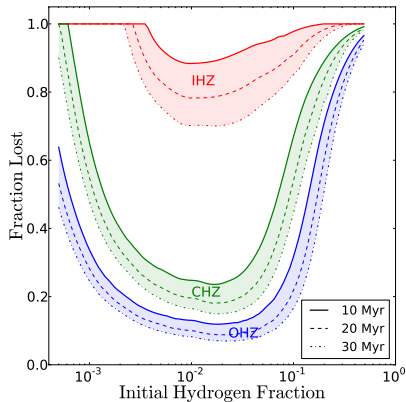


The Effect of Eccentricity

$e = 0$



$e = 0.7$









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Conclusions

-  HECs can form from mini-Neptunes with $f_{H_0} \gtrsim 0.25$ **scattered** into the HZ of $M \lesssim 0.2M_{\odot}$ M dwarfs
-  HECs more likely for **high** f_{H_0}
-  HECs form **early** ($t \lesssim 50$ Myr)
-  Hydrodynamic escape, Roche lobe overflow, tidal evolution and thermal evolution all play a **critical** role in forming HECs
-  This process may be the primary mechanism for the formation of **habitable planets** around M dwarfs.
-  HECs may be observed in the **next few years**





Thank you



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