

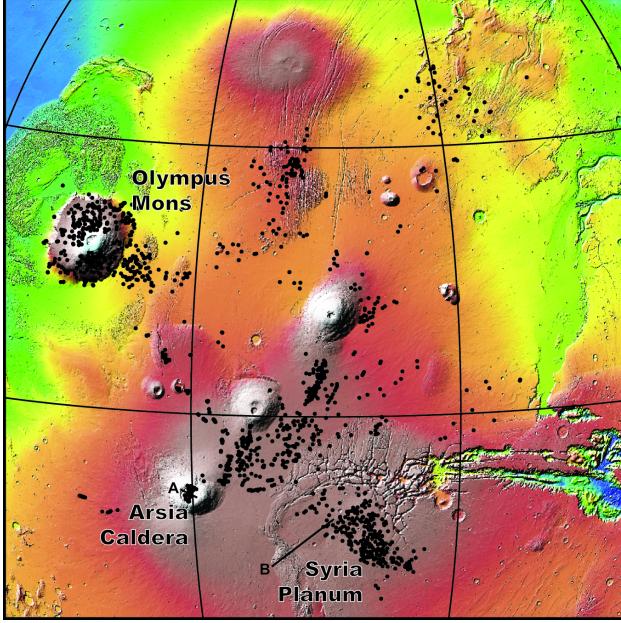
# Recurrence rate and magma effusion rate for the latest volcanism on Arsia Mons, Mars

## 1 Introduction

Greeley and Schneid [1991] produced one of the first extrusive magma flux estimates for the surface of Mars and used terrestrial intrusive/extrusive ratios to calculate that  $6.5 \cdot 10^8 \text{ km}^3$  of magma has been generated on Mars in the past 3.8 Ga. For the most recent 500 Ma, magma production was observed to wane, and only  $2.11 \cdot 10^6 \text{ km}^3$  of magma was modeled to have erupted [Greeley and Schneid, 1991]. This global extrusive magma flux of  $0.004 \text{ km}^3/\text{yr}$  ( $0.13 \text{ m}^3/\text{s}$ ) remains one of few such estimates of magma production. Constraining past and recent magma production and extrusion rates, however, is of vital importance in understanding evolution of the Martian climate [Mouginis-Mark and Rowland, 2008, Halevy and Head III, 2014, e.g.], lithosphere and mantle [Grott et al., 2013], surface [Wilson and Head, 1994], and the ability of Mars to sustain biotic or pre-biotic material over time [Scanlon et al., 2015b].

The large volcanic edifices of the Tharsis region have been given a time-averaged magma flux estimate of  $0.05 \text{ m}^3/\text{s}$ , with a factor of 3 uncertainty, for an active construction period of 1 Gyr (again with a factor of 3 uncertainty) by Wilson et al. [2001]. Wilson et al. [2001] further constrained periodic flux under these volcanoes by assuming each of the many summit calderas was formed in association with one stable magma body at depth. By calculating the necessary flux to achieve the magma bodies that could form such calderas, Wilson et al. [2001] found that the magma delivery rate to the volcanoes had to persist at  $1\text{-}10 \text{ m}^3/\text{s}$  for hundreds of thousands to millions of years, followed by orders of magnitude longer periods of quiescence before new large magma bodies could be emplaced. For example, assuming a magma chamber size of  $50,000 \text{ km}^3$  under the Arsia Mons caldera [Wilson et al., 2001], it would take a  $10 \text{ m}^3/\text{s}$  magma flux 140 kyr to fill the magma chamber, representing  $\sim 3\%$  of Arsia's total volume in just 0.01% of a 1 Gyr constructional period.

We seek to estimate both the recurrence rate of volcanism and the extrusive magma delivery rate for the most recent volcanic unit on Arsia Mons—a patchwork of lava flows and 29 associated volcanic vents within the volcano caldera. To constrain these values, absolute ages and associated age uncertainty of each flow have been modeled using the size-frequency distribution of observed craters, relative ages between flows have been determined using superposition relationships, and volumes have been modeled using Mars Orbiter Laser Altimeter (MOLA) data [Smith et al., 2003]. Crater-retention age models and stratigraphic relations are integrated in a Volcanic Event Recurrence Rate Model (VERRM) to better



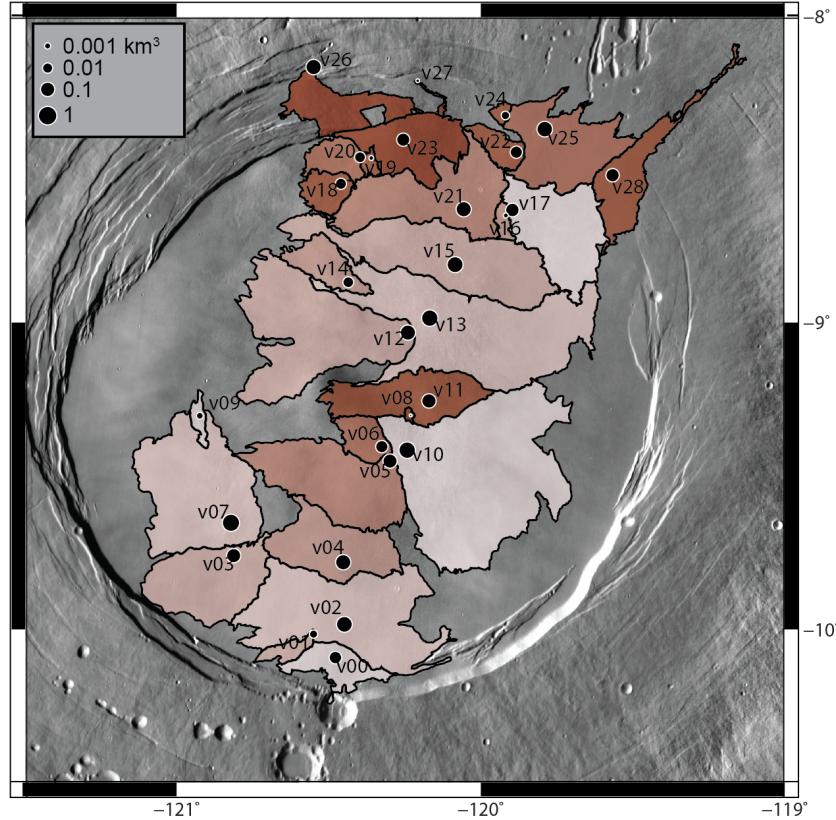
**Figure 1:** The Tharsis Volcanic Province of Mars. Each black dot represents one small volcanic vent in a Tharsis-wide catalog. Within Arsia Mons (bottom left), 29 vents have been identified. The color relief basemap is MOLA topography.

characterize event age uncertainty and estimate recurrence rate throughout the period of time during which these lavas were emplaced.

## 2 Geologic Background of Arsia Mons

Arsia Mons is a major shield volcano on Mars and a member of the Tharsis Montes (Figure 1). With a diameter of over 300 km and slopes of  $5^\circ$  [Plescia, 2004], the surface of Arsia contains lava flows, which served as the primary construction material of the shield [Mouginis-Mark and Rowland, 2008], prodigious ash deposits [Mouginis-Mark, 2002], and glacial deposits [Head and Marchant, 2003] emplaced under both cold- and warm-based glacial conditions [Scanlon et al., 2015b]. At the summit of Arsia is a single collapse caldera measuring  $\sim 4000$   $\text{km}^3$  in volume [Wilson et al., 2001]. Within this 110 km wide caldera, a linear cluster of secondary shield volcanoes comprise one of the youngest geologic units in the Arsia region [Carr et al., 1977, Scott and Zimbelman, 1995]. No craters larger than 1 km exist within the caldera and several detailed crater retention studies with different image datasets have independently produced 130 Ma as a single age estimate of the entire caldera floor [Neukum et al., 2004, Werner, 2009, Robbins et al., 2011].

Through Mariner 9 and Viking Orbiter images of the Tharsis region, Crumpler and Aubele [1978] determined that of the three Tharsis Montes, Arsia Mons is the most structurally evolved shield volcano. This conclusion was backed up by Bleacher et al. [2007a] with extensive mapping using HRSC and THEMIS. Bleacher et al. [2007a] suggested that the northward trending structural complexity of the Tharsis Montes might indicate a migrating magma source along the axis of the three volcanoes, similar to a Tharsis plume model



**Figure 2:** Mapped lava flows in the vent field. Circles are placed at the 29 vent locations and are sized according to volume. Lighter colored flows are higher stratigraphically than darker colored flows (See Figure 3).

by Mège and Masson [1996]. If such a migrating magma source did exist, then magma production at Arsia would have waned, decreasing the amount of melt available for continued summit volcanism.

On the flank of Arsia Mons, Mouginis-Mark and Rowland [2008] identified >1000 layered units in a graben, which were interpreted to be lava flows. Using MOLA data, they were able to estimate the height of the graben wall, enabling the estimation of layer thicknesses between 10-80 m, with most being >30 m. As no unique and laterally extensive layers were observed in the stack of <2 km wide layers, Mouginis-Mark and Rowland [2008] concluded that no major glacial events were emplaced between the deposition of these layers, perhaps indicating relatively rapid emplacement of 885 m of lava. By instead assuming constant activity of Arsia Mons for either 2 or 3 billion years, Mouginis-Mark and Rowland [2008] estimated that the stack could have been emplaced over 290 or 435 million years, respectively.

## 2.1 Other recent volcanism in the Tharsis Region

Outside of Arsia, several volcanic events and trends in the Amazonian Period (<3 Ga) have been identified. On Olympus Mons, the latest flank lava flows appear to have transitioned from long tube-fed flows to shorter, channel-fed flows, suggesting that magma flux per flow

emplacement event waned over the late Amazonian [Bleacher et al., 2007b]. Bleacher et al. [2007b] concluded that this might be evidence of a larger waning of activity at Olympus and might also signal a transition from a deep mantle source of volcanism to shallower sources, based on long-term volcanic patterns observed at the Hawaiian volcanic chain. These observations might be consistent with the Wilson et al. [2001] model that most of the history of these volcanoes is spent dormant after hyper-active edifice building episodes wane and end.

Recent volcanism elsewhere in Tharsis has occurred in a more distributed fashion [Hauber et al., 2011], including a lava field filling the southeast “moat” of Olympus Mons [Chadwick et al., 2015]. Chadwick et al. [2015] identified lava flows at the base of Olympus emplaced between 64-210 Ma, similar in time to previous age estimates of the study area in this paper ( $\sim$ 130 Ma). The emplacement direction of these lavas are systematically offset from the current downward direction, due to recent subsidence of Olympus Mons. This might provide evidence contrar to he findings of Bleacher et al. [2007b], as Chadwick et al. [2015] estimate that the magma required to be injected into Olympus Mons the past 200 Myr in order to cause this subsidence might be  $10^5$ - $10^6$  km $^3$ . This would suggest that Olympus is not yet extinct.

## 3 Methods

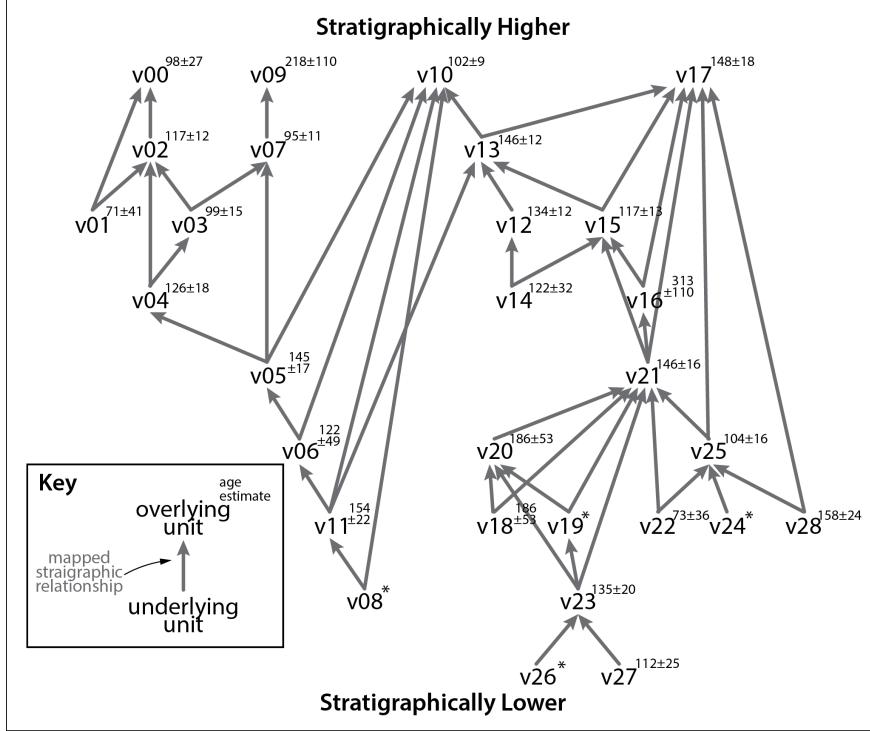
### 3.1 Unit and stratigraphic mapping

The 29 mapped volcanic vents within the caldera each have lava flows emanating from them that form positive topographic features over the surrounding terrain. Flows corresponding to each vent have been mapped in ArcGIS 10.2 with georeferenced Context Imager (CTX) photographs [Malin et al., 2007] serving as a 6-m resolution basemap (Figure 2). Flows are mapped in association with an observed vent where flows can be unambiguously traced directly back to the vent using flow features. Some lava flows on the eastern and western margins of the caldera appear to flow away from the caldera center and might have been created during an event that formed any of the observed vents; however, because they are covered in subsequent flows and are separated from their parent vent by at least one flow front, they cannot be traced to a vent and are not included in our catalog.

Mapped flows which abut each other have an inherent superposition relationship. Using available CTX images, these relationships are documented for all neighboring flows by identifying features such as 1) diverted flows around pre-existing topography, 2) infilling of graben or volcanic vents, and 3) continuous flow features suddenly vanishing under overlying flows. Superposition relationships are graphed according to stratigraphic height in Figure 3.

### 3.2 Crater retention age modeling

Mapped units within the caldera (29 lava flow units and 1 “undifferentiated lavas” unit) have been assigned model ages based on the distribution of impact craters within their boundaries. Robbins et al. [2011] previously found that, in the Arsia Mons caldera, crater frequency decreases for craters with diameters (D) of  $\leq$  93 m, due to dust cover. Robbins et al. [2011] also hypothesized that a background population of secondary craters in this area might contaminate the crater distribution at D < 130 m. We have counted all craters

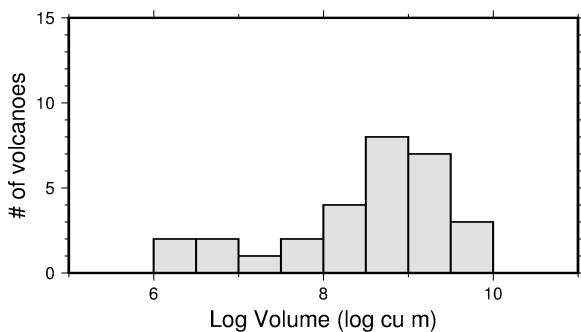


**Figure 3:** A directed graph of stratigraphic relationships within the Arsia Mons Caldera. Graphically higher flows (e.g. v00) are higher stratigraphically and directly overlay flows that have connecting lines to them. Each line is one mapped stratigraphic relationship connecting an overlying flow and an adjacent underlying flow. Age estimates from craterstats2 are given in superscript; \* symbols indicate no age was estimated due to a lack of craters with  $D \geq 100$  m.

in the caldera with  $D \geq 100$  m to avoid both of these systematic biases. Four lava flows have an exposed area smaller than  $15 \text{ km}^2$  and an insignificant number of craters larger than 100 m were found on their surfaces (10 or fewer), so no crater age date was determined. The diameter and location of craters were mapped in ArcGIS 10.2 using a basemap of CTX images.

Modeled ages based on a impact production function, which is a model of the relative production rate of different impact crater sizes, and a chronology function, which models the change of impact frequency with time. Ages were modeled in the Craterstats2 software [Michael, 2013]. For each mapped unit, craters are separated by diameter into bins of minimum diameter  $2d$  km where  $d$  increments by 0.5 between each bin, following the Hartmann 2004 iteration Production Function [Hartmann, 2005]. Ages are modeled based on this Production Function and the Michael [2013] Chronology Function. Uncertainty of crater frequency in each bin is defined as  $\sqrt{N}/S$ , where  $N$  is the number of craters and  $S$  is the area of the mapped unit. The cumulative crater-size frequency distribution (and associated uncertainty) is assembled from each  $2d$  km bin and is used to model a best fit age (see Figure 8. Ages and uncertainties for each mapped unit are reported in Ma).

Histogram of 'Minimum' Volumes at Arsia



**Figure 4:** Minimum estimated volumes of each lava flow using an interpolated mesh as a theoretical subsurface. All lavas are estimated to be between  $10^{-3}$ - $1 \text{ km}^3$ , with most volumes centering around  $1 \text{ km}^3$ .

### 3.3 Volume estimation

The thicknesses of the mapped flows are unknown, which makes estimating flow volume a challenge. Thicknesses of lava flows on the flanks of Arsia Mons have been measured to be 10-80 m [Mouginis-Mark and Rowland, 2008]. These thicknesses are in agreement with other studies of lava flow thickness, including flows in Elysium Mons with thicknesses of 7-35 m [Pasckert et al., 2012], flows on Ascreaus Mons between 24-88 m [Hiesinger et al., 2007], and 37 m and 50 m thick flows in the Elysium region and on Pavonis Mons [Glaze et al., 2003, Baloga et al., 2003]. If the range reported by Mouginis-Mark and Rowland [2008] is taken to be upper and lower thickness bounds for each lava flow, a volume range can be calculated for each flow by multiplying these bounds by the exposed area.

An alternative volume estimation method would be to derive flow thickness using a subsurface model of each mapped unit. In this method, elevation values are assigned to the mapped perimeter of each flow from the gridded-MOLA topographic dataset. The subsurface of the flow is modeled with a triangular irregular network (TIN) where mesh faces are generated with the Delaunay triangulation of vertices along the perimeter. The modeled subsurface is then subtracted from the MOLA grid, producing a thickness map which is integrated to estimate flow volume. This estimate should be considered to be a minimum estimate of flow volume, because the modeled sub-flow surface tightly connects the lava flow margins, while in real life lava flows invert topography and should likely have a deeper, more concave upward subsurface. Though these volumes are likely underestimates, they are modeled using the topography of each edifice instead of using nearby flow measurements of different lavas as is described above.

### 3.4 Volcanic Event Recurrence Rate Model (VERRM)

Together, stratigraphic information and crater retention age estimates can be consolidated to improve age uncertainty estimation for volcanic events. This is especially applicable to recent volcanic landforms on Mars as crater-based dates alone might be biased due to crater burial, local topography, or secondary crater background populations [Robbins et al., 2011, Platz and Michael, 2011].

To accomplish the task of constraining modeled age estimates with stratigraphy, and ultimately to describe the repose interval of volcanic events in the region, we have devised an algorithm, which we call the Volcanic Event Recurrence Rate Model (VERRM). VERRM implements a Monte Carlo algorithm that assigns a potential age to each volcanic event by defining an event age distribution function,  $A$ , with an event's modeled crater retention rate age and age uncertainty, as reported by Craterstats2. The age distribution function for each event is modeled as a Gaussian distribution, which is a suitable for the events in this study as they are relatively recent and the crater impact rate during the time of their emplacement is thought to be constant [Vaucher et al., 2009]. The initial age function is given as  $A_i(\mu_i, \sigma_i^2)$  where  $\mu$  is the estimated age determined by crater retention and  $\sigma$  is the uncertainty of the estimated age of cataloged event  $i$ . If no crater retention age estimate is given (e.g. for smaller flows with an insignificant number of impact craters), the initial age function is assumed to be a uniform probability function bounded by a minimum and maximum possible age described next.

VERRM then constrains  $A$  with a binary stratigraphy function, with possible ages having a value of 1 and ages outside an acceptable age range having a value of 0. Possible ages are defined by previously dated events in the VERRM simulation; stratigraphically higher events connected to the event at hand give a minimum age of the stratigraphy function, while lower events give a maximum age. If no stratigraphically-related events have been dated, the minimum age bound is set to the present, and the maximum age bound is set to three times older than the oldest crater retention rate modeled age in this field (330 Ma), or 1 Ga.

The normalized product of the the Gaussian age function and the binary stratigraphy function gives an age distribution function which does not violate stratigraphy but is informed by crater retention age estimates. This function is sampled to date an event and the process repeats for the next event. By repeating this process 10,000 times, the potential age ranges for each event is determined.

**Recurrence rate calculation** Volcanic recurrence rate RR at any time  $t$  is modeled for each MC solution as

$$RR(t) = \frac{n - 1}{\max\{\mathbf{T}\} - \min\{\mathbf{T}\}} \quad (1)$$

where  $n$  is an even number and  $\mathbf{T}$  is the set of ages of  $n$  events temporally surrounding  $t$ . Events whose ages are elements of set  $\mathbf{T}$  are identified as the closest  $n/2$  events either before or after  $t$ . The even number  $n$  is essentially a local window in time, and  $n - 1$  is the number of recurring events between  $\max\{\mathbf{T}\} - \min\{\mathbf{T}\}$  after the earliest event occurrence. Note that all times  $t$  laying between two temporally consecutive events will have the same set  $\mathbf{T}$  and will therefore have the same calculated recurrence rate.

The set  $\mathbf{T}$  is also modified for early and late times in the MC solution. For times before the earliest event in an individual simulation (i.e. where  $t$  is greater than the time of the oldest age assigned to an event),  $\mathbf{T}$  includes the earliest  $n - 1$  events and time  $t$ . When  $t$  falls in time between the earliest event and the  $n/2^{\text{th}}$  event to occur in the simulation,  $\mathbf{T}$  includes the first  $n$  events. Likewise, for later (i.e. more recent) times, where  $t$  is more recent than the earliest event,  $\mathbf{T}$  includes both  $t$  and the most recent  $n - 1$  events. Where  $t$  falls in time between the most recent event and the  $n/2^{\text{th}}$  most recent event,  $\mathbf{T}$  is the set of the most

recent  $n$  events. For the cases where  $t$  is at the same time as the earliest or latest events,  $\mathbf{T}$  will be the set of times of the  $n$  earliest or latest events, respectively.

While recurrence rate for any one MC simulation might be bad, the range of recurrence rates for all simulations can provide a confidence envelope of recurrence rate with respect to time. The recurrence rate for the field is defined as the median recurrence rate of all simulations over time, with a 95% confidence envelope serving as the rate's uncertainty.

**Volume flux calculation** Volume flux is also modeled with the VERRM output. Each event is assigned a volume (discussed above in section 3.3) for the lava flow it produced. Volume flux for the field is defined as the volume of erupted lava per million year time bin. A 95% confidence envelope and median flux through time are obtained using all MC solutions.

## 4 Results

The 29 lava flows in our catalog are mapped to cover 6700 km<sup>2</sup>, ~70%, of the caldera, representing the majority of surface lavas within the caldera walls (Figure 2). Each flow has at least one neighboring flow, which enables a stratigraphic web to be produced that connects all volcanic events to other events 3. Age estimates modeled from crater retention ranged from 71-313 Ma, though the youngest and oldest ages are not observed at the top or bottom of the stratigraphic web. This indicates that the crater ages might not be well constrained for small and relatively young surface units on Mars, but that ages for these units might be better constrained with stratigraphic information. These data are detailed for each edifice in Table 1.

Lava flow volume estimates using the subsurface modeling method range from  $1.6 \cdot 10^{-3}$  to 5 km<sup>3</sup>, with an average of 1.1 km<sup>3</sup> per flow (Figure 4). The total volume using this method, again assumed to be a minimum volume estimate, is 31 km<sup>3</sup>. Average flow thickness for the 29 flows is 4.6 m, which is thinner than the thinnest flows identified by Mouginis-Mark and Rowland [2008] by a factor of two. Individual volumes are reported in Table 2. We also report the volumes of each flow assuming an average thickness of 10 and 80 m, which might be more reasonable volume bounds. Assuming these two thicknesses gives the flow field a total modeled volume of 67 or 540 km<sup>3</sup>, respectively.

## 5 Discussion

### 5.1 Comparisons to other studies

Modeled ages of these flows with our crater counts lay between 70-400 Ma, with uncertainties reported by craterstats2 to be between 10-100 Myr. Our ages confirm modeled ages produced by other authors, where we find the crater-derived model age of the entire caldera to be 130 Ma, similar to Neukum et al. [2004], Werner [2009] and [Robbins et al., 2011]. Robbins et al. [2011] also used crater retention rates (for D93 m) to date a single endogenous crater located within the caldera at 9.70S, 239.18E as having an age of  $97 \pm 49$  Ma. We independently date the lava flow associated with this vent to have an age of  $99 \pm 15$  Ma (Figure 8).

**Table 1:** Lava Flow Age Information

Flow	Vent Long.	Location Lat.	Modeled Age (Ma)	Overlying Flows	Underlying Flows
v00	-120.48°	-10.09°	98.4±27	—	02,01
v01	-120.55	-10.02	71.4±41	00,02	—
v02	-120.45	-9.98	117±12	00	01,03,04
v03	-120.81	-9.76	99.1±15	02,07	04
v04	-120.45	-9.78	126±18	02,03	05
v05	-120.30	-9.45	145±17	04,07,10	06
v06	-120.33	-9.41	122±49	05,10	11
v07	-120.82	-9.65	94.7±11	09	03,05
v08	-120.23	-9.30	—	11,10	—
v09	-120.92	-9.30	218±110	—	07
v10	-120.24	-9.42	102±9.3	—	05,06,08,11,13
v11	-120.17	-9.26	154±22	06,10,13	08
v12	-120.24	-9.03	134±12	13	14
v13	-120.17	-8.99	146±12	10,17	12,15,11
v14	-120.44	-8.87	122±32	12,15	—
v15	-120.08	-8.81	117±13	13,17	14,16,21
v16	-119.92	-8.65	313±110	15,17	21
v17	-119.90	-8.63	148±18	—	13,15,16,21,25,28
v18	-120.46	-8.55	186±53	20,21	—
v19	-120.36	-8.46	—	20,21	23
v20	-120.40	-8.46	186±53	21	18,19,23
v21	-120.06	-8.63	146±16	15,16,17	20,18,19,23,22,25
v22	-119.88	-8.44	73.2±36	21,25	—
v23	-120.25	-8.40	135±20	19,20,21	26,27
v24	-119.92	-8.32	—	25	—
v25	-119.79	-8.37	104±16	17,21	22,24,28
v26	-120.21	-8.21	—	23	—
v27	-120.55	-8.17	112±25	23	—
v28	-119.57	-8.52	158±24	17,25	—

Our initial VERRM results suggest that the Arsia field might have been active for 230 Myr, ending about 70 Mya. If  $280 \text{ km}^3$  of basalt was emplaced as lavas during this time, the time-averaged volume flux of the field would have been  $1,200 \text{ m}^3 \text{ a}^{-1}$ . ( $3.6 \cdot 10^{-4} \text{ m}^3 \text{ s}^{-1}$ ). This is two orders of magnitude less active than the magma flux estimated for Central Elysium Planitia, calculated by Vaucher et al. [2009] to be  $1.4 \cdot 10^{-2}$  to  $1.8 \cdot 10^{-2} \text{ m}^3/\text{s}$  over the most recent 234 Myr, through similar volume estimates of lava flows and a crater retention rate study. Our estimate is also 5 orders of magnitude lower than the average magma flux ( $30 \text{ m}^3 \text{ s}^{-1}$ ) Wilson et al. [2001] calculated would be needed to charge the most recent magma chamber under Arsia Mons. If we employ the same 8.5:1 intrusive/extrusive ratio used by Greeley and Schneid [1991], the total magma production at depth during the emplacement

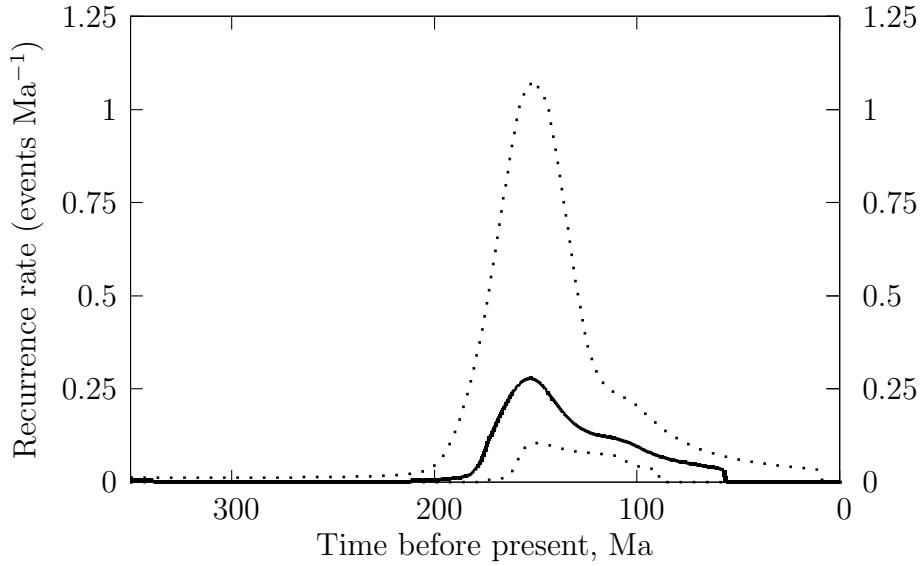
**Table 2:** Lava Flow Morphometry

Flow	Vent Location		Area (km <sup>2</sup> )	Volume Estimates* (km <sup>3</sup> )		
	Long.	Lat.		subs.	model	10 m
v00	-120.48°	-10.09°	102.27	0.37	1.0	8.2
v01	-120.55	-10.02	31.21	0.032	0.31	2.5
v02	-120.45	-9.98	538.86	2.3	5.4	43
v03	-120.81	-9.76	283.32	0.91	2.8	23
v04	-120.45	-9.78	242.29	1.6	2.4	19
v05	-120.30	-9.45	353.98	0.89	3.5	28
v06	-120.33	-9.41	69.40	0.31	0.69	5.6
v07	-120.82	-9.65	479.32	5.0	4.8	38
v08	-120.23	-9.30	4.00	0.0017	0.04	0.32
v09	-120.92	-9.30	21.08	0.013	0.21	1.7
v10	-120.24	-9.42	858.69	3.2	8.6	69
v11	-120.17	-9.26	228.09	1.0	2.3	18
v12	-120.24	-9.03	528.14	1.6	5.3	42
v13	-120.17	-8.99	679.49	3.2	6.8	54
v14	-120.44	-8.87	96.78	0.14	0.97	7.7
v15	-120.08	-8.81	460.75	2.5	4.6	37
v16	-119.92	-8.65	18.11	0.0053	0.18	1.4
v17	-119.90	-8.63	300.03	0.64	3.0	24
v18	-120.46	-8.55	64.29	0.16	0.64	5.1
v19	-120.36	-8.46	13.12	0.0047	0.13	1.0
v20	-120.40	-8.46	70.51	0.14	0.71	5.6
v21	-120.06	-8.63	354.97	1.6	3.5	28
v22	-119.88	-8.44	53.20	0.31	0.53	4.3
v23	-120.25	-8.40	204.71	0.53	2.0	16
v24	-119.92	-8.32	11.70	0.037	0.12	0.94
v25	-119.79	-8.37	281.19	2.0	2.8	22
v26	-120.21	-8.21	5.61	0.0026	0.056	0.45
v27	-120.55	-8.17	158.16	1.8	1.6	13
v28	-119.57	-8.52	176.53	0.58	1.8	14

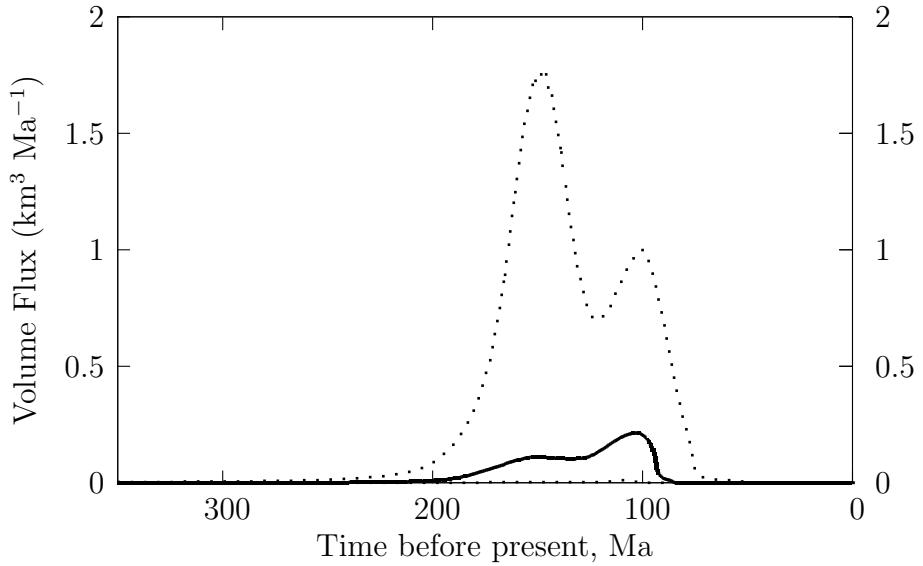
\*The three volume estimates are found using 1) a modeled subsurface, 2) assuming an average thickness of 10 m and 3) 80 m.

of the caldera flows would be  $3.4 \cdot 10^{-3} \text{ m}^3 \text{ s}^{-1}$ ). This is still much lower than the estimated flux necessary to sustain a magma chamber, but is only one order of magnitude smaller than the average flux needed to build Arsia Mons in 1 Gyr,  $0.05 \text{ m}^3 \text{ s}^{-1}$  [Wilson et al., 2001].

Time-averaged recurrence can be estimated by dividing the total elapsed time of volcanic activity by number of events. For instance, Richardson et al. [2013] identified 263 monogenetic volcanic vents within Syria Planum, which were interpreted to be emplaced from 3.6-2.9 Ga, or 700 Myr. If volcanism were constant in that area, a new volcanic vent would

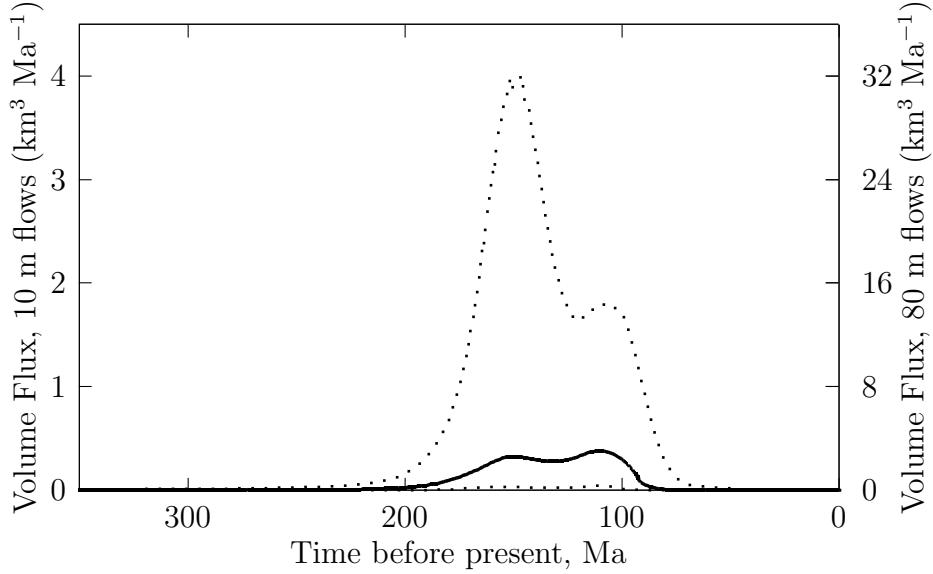


**Figure 5:** Subsurface Volume Method



**Figure 6:** Subsurface Volume Method

have been formed every 2.7 Myr. In a graben on the northwest flank of Arsia, Mouginis-Mark and Rowland [2008] mapped >1000 lava flows and estimated construction rates of 290 and 435 million years, based on the time to build all of Arsia Mons. This would correspond to a recurrence of at least one episode of lava emplacement every 290 or 440 kyr. A time-averaged recurrence for our 29 vents, created over 230 Myr, would be one event every 8 Myr. This alone would imply that the latest volcanic activity on Arsia Mons was much closer in style to the volcanism on Syria Planum than during the main constructional phases of Arsia Mons.



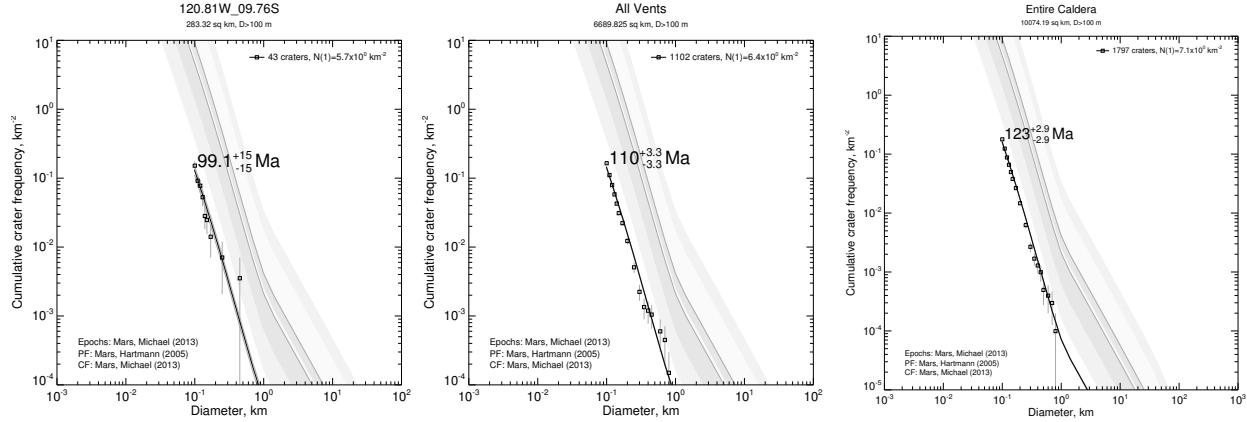
**Figure 7:** Subsurface Volume Method

## 5.2 Effects on tropical mountain glaciers on Arsia Mons

In the past decade, studies have interpreted fan-shaped deposits on the western flanks of the Tharsis Montes to be recent glacial deposits, due to the presence of fresh moraines and possible stranded ice blocks, analogous to kettles on Earth [Shean et al., 2007, Kadish et al., 2014, Scanlon et al., 2015a]. The material on these broad deposits have been dated by Kadish et al. [2014] to have been emplaced around 200 Ma. Scanlon et al. [2015b] identified fan-shaped deposits on the western flank of Arsia Mons which contain evidence of basal melting in clear association with sub-glacial volcanic eruptions.

Recent analysis of smooth facies deposits to the northwest of the Arsia summit has provided evidence that tropical mountain glaciers are, in fact, extant and covered in ash [Scanlon et al., 2015a]. The penetration depth of viscously relaxed ring-mold craters on these deposits indicates a maximum material blanket thickness of less than 230 m over tens to hundreds of meters of ice or ice-rich material [Head and Weiss, 2014]. A large portion of this insulating material might in fact be volcanic ash [Wilson and Head, 1994, Mouginis-Mark, 2002].

The presence of volcanic ash on remaining tropical mountain glaciers on the flank of Arsia might be a result of the most recent volcanism in the Arsia caldera. If this is the case, our volume estimates would likely be severely underestimated, as a large portion of erupted material would have been transported away from the vent as tephra. However, as Kadish et al. [2014] estimated the resurfacing age of the fan-shaped deposits to be about 200 Ma, roughly our interpreted age of onset of effusive activity within the caldera, it is possible that the ash on the flanks and the lavas in the caldera represent a transition from explosive to effusive volcanism at Arsia Mons. The provenance of the ash might be buried by the recent lavas or might be the “parasitic calderas” observed by Crumpler et al. [1996], which form a rift of the south and north caldera walls of Arsia, in line with the shield volcanoes in our catalog.



**Figure 8:** Cumulative crater frequency distributions for (left) one lava flow, (center) all mapped lava flows, (right) the entire caldera). The example flow, v03 in our database, had been previously estimated to be  $97 \pm 49$  Ma from crater counting performed on a HiRISE image [Robbins et al., 2011]. Our age of  $99 \pm 15$  Ma agrees with this previous finding. The age of the entire caldera also generally agrees with previous estimates of 130 Ma. The mapped vents plot younger than the entire caldera, which is expected as they are the most recent resurfacing events on the caldera.

### 5.3 Transition of volcanic style and waning of Arsia volcanism

From our findings and evidence of prodigious amounts of ash to the west [Mouginis-Mark, 2002], we interpret the intra-caldera lava flows to represent the waning of activity at Arsia Mons, which effectively ended  $\sim 70$  million years ago. Preceding the formation of the large basaltic caldera, magmatic activity would have mainly been related to a large magma chamber, which was active at a heightened rate [Wilson et al., 2001]. Dikes ascending below this chamber are expected to propagate towards the chamber and ultimately assimilate with it [Karlstrom et al., 2009], effectively suppressing monogenetic volcanism directly above the chamber during its molten history. Instead, large events related to the central chamber would have occurred, perhaps supplying the provenance of ash now mantling ice-rich material on Arsia’s west flanks.

The relatively small volumes and distributed nature of the lavas within the caldera, as well as the caldera itself, are evidence that the magma chamber has cooled and that dikes are able to ascend individually through the center of the caldera. The magma supply rate to the base of Arsia must therefore have waned to below  $\sim 12 \text{ m}^3 \text{ s}^{-1}$  (plus or minus a factor of three) [Wilson et al., 2001]. The long-term average surface flux of distributed volcanism in this study,  $3.6 \cdot 10^{-4} \text{ m}^3 \text{ s}^{-1}$ , is five orders of magnitude below this minimum supply rate.

The confidence envelope for volcanic recurrence rate for times older than the peak of volcanic activity in this study (XXX Ma) is largely unconstrained; while the envelope suggests that recurrence among the 29 volcanoes waxes during this time, it cannot be used to imply that volcanic activity in general waxed during this time. It is beyond likely that below the 29 lava flows, other flows and volcanic vents are buried. The waxing trend seen in our results might be erroneous because of a lack of data on buried events. However, it is likely that the

waning trend observed since XXX Ma is real.

## 6 Conclusions

The 29 lava flows mapped in the caldera create a volcanic field that was active since at least 230 Ma and became inactive at around 70 Ma. These flows represent 30-400 km<sup>3</sup>, giving a surficial magma flux rate of  $3.6 \cdot 10^{-4}$  m<sup>3</sup> s<sup>-1</sup>, averaged over the entire history of the 29 events.

Our main conclusion is that the recurrence rate of intra-caldera volcanism on Arsia Mons has waned since XXX Ma and that this waning, as well as the presence of distributed lava flows in the caldera alone, represent the tail-end of an episode of activity in Arsia Mons that formed the large basaltic caldera and a corresponding magma chamber hypothesized by Wilson et al. [2001].

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## 7 Appendix: cumulative crater charts

