

# Eclipse

en.wikipedia.org/wiki/Eclipse

This article is about astronomical eclipses. For other uses, see [Eclipse \(disambiguation\)](#).

"Total eclipse" redirects here. For other uses, see [Total Eclipse \(disambiguation\)](#).



Totality during the 1999 solar eclipse. [Solar prominences](#) can be seen along the limb (in red) as well as extensive [coronal](#) filaments.

An **eclipse** is an [astronomical event](#) that occurs when an [astronomical object](#) is temporarily obscured, either by passing into the shadow of another body or by having another body pass between it and the viewer. An eclipse occurs during a [syzygy](#).<sup>[1]</sup>

The term eclipse is most often used to describe either a [solar eclipse](#), when the Moon's shadow crosses the Earth's surface, or a [lunar eclipse](#), when the Moon moves into the Earth's shadow. However, it can also refer to such events beyond the Earth–Moon system: for example, a planet moving into the shadow cast by one of its moons, a moon passing into the shadow cast by its host planet, or a moon passing into the shadow of another moon. A [binary star](#) system can also produce eclipses if the plane of the orbit of its constituent stars intersects the observer's position.

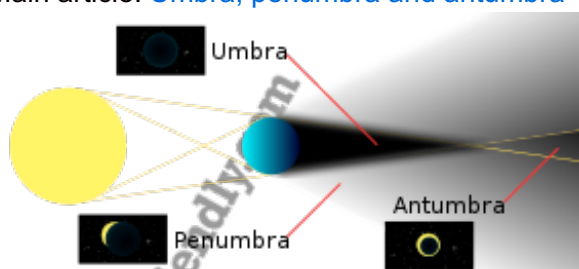
## Contents

## Etymology

The term is derived from the [ancient Greek](#) noun ἐκλειψις (*ékleipsis*), which means "the abandonment", "the downfall", or "the darkening of a heavenly body", which is derived from the verb ἐκλείπω (*ekleípō*) which means "to abandon", "to darken", or "to cease to exist",<sup>[2]</sup> a combination of prefix ἐκ- (*ek-*), from preposition ἐκ (*ek*), "out," and of verb λείπω (*leípō*), "to be absent".<sup>[3][4]</sup>

## Umbra, penumbra and antumbra

Main article: [Umbra, penumbra and antumbra](#)

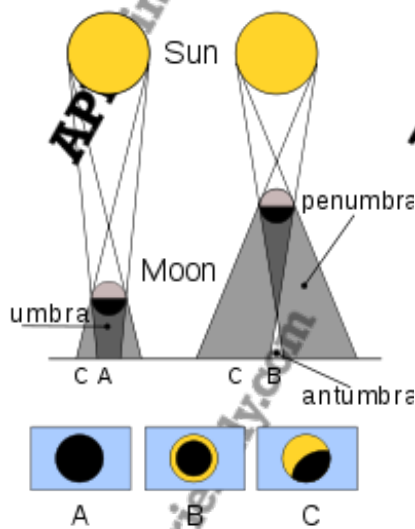


Umbra, penumbra and antumbra cast by an opaque object occulting a larger light source

For any two objects in space, a line can be extended from the first through the second. The latter object will block some amount of light being emitted by the former, creating a region of shadow around the axis of the line. Typically these objects are moving with respect to each other and their surroundings, so the resulting shadow will sweep through a region of space, only passing through any particular location in the region for a fixed interval of time. As viewed from such a location, this shadowing event is known as an eclipse.<sup>[5]</sup>

Typically the cross-section of the objects involved in an astronomical eclipse are roughly disk shaped.<sup>[5]</sup> The region of an object's shadow during an eclipse is divided into three parts:<sup>[6]</sup>

- The **umbra**, within which the object completely covers the light source. For the **Sun**, this light source is the photosphere.
- The **antumbra**, extending beyond the tip of the umbra, within which the object is completely in front of the light source but too small to completely cover it.
- The **penumbra**, within which the object is only partially in front of the light source.



Sun-moon configurations that produce a total (A), annular (B), and partial (C) solar eclipse

A **total eclipse** occurs when the observer is within the umbra, an annular eclipse when the observer is within the antumbra, and a partial eclipse when the observer is within the penumbra. During a lunar eclipse only the umbra and penumbra are applicable. This is because Earth's apparent diameter from the viewpoint of the Moon is nearly four times that of the Sun. The same terms may be used analogously in describing other eclipses, e.g., the antumbra of Deimos crossing Mars, or Phobos entering Mars's penumbra.

The *first contact* occurs when the eclipsing object's disc first starts to impinge on the light source; *second contact* is when the disc moves completely within the light source; *third contact* when it starts to move out of the light; and *fourth or last contact* when it finally leaves the light source's disc entirely.

For spherical bodies, when the occulting object is smaller than the star, the length ( $L$ ) of the umbra's cone-shaped shadow is given by:

$$L = \frac{r \cdot R_o}{R_s - R_o}$$

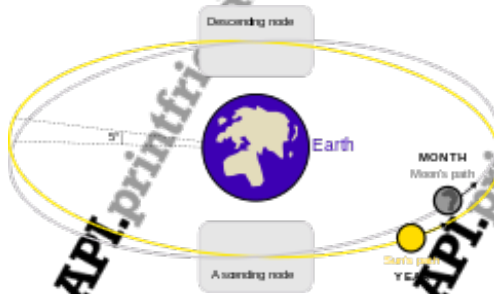
where  $R_s$  is the radius of the star,  $R_o$  is the occulting object's radius, and  $r$  is the distance from the star to the occulting object. For **Earth**, on average  $L$  is equal to  $1.384 \times 10^6$  km, which is much larger than the Moon's **semimajor axis** of  $3.844 \times 10^5$  km. Hence the umbral cone of the Earth can completely envelop the Moon during a

On Earth, the shadow cast during an eclipse moves very approximately at 1 km per sec. This depends on the location of the shadow on the Earth and the angle in which it is moving.

<http://www.sciforums.com/threads/speed-of-eclipse-shadow.53722/>

Main article: [Eclipse cycle](#)

## Earth–Moon System



An eclipse involving the Sun, Earth, and Moon can occur only when they are nearly in a straight line, allowing one to be hidden behind another, viewed from the third. Because the **orbital plane of the Moon** is tilted with respect to the orbital plane of the Earth (the **ecliptic**), eclipses can occur only when the Moon is close to the intersection of these two planes (the **nodes**). The Sun, Earth and nodes are aligned twice a year (during an **eclipse season**), and eclipses can occur during a period of about two months around these times. There can be from four to seven eclipses in a calendar year, which repeat according to various **eclipse cycles**, such as a **saros**.

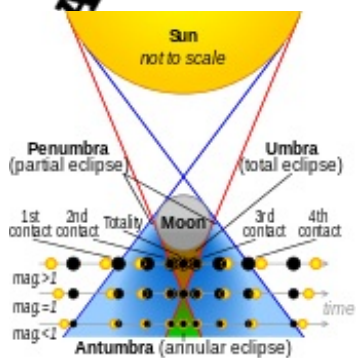
- four (penumbral) lunar and three solar eclipses: 1908, 2038.
- four solar and three lunar eclipses: 1917, 1973, 2094.
- five solar and two lunar eclipses: 1934.

- 1591, 1656, 1787, 1805, 1917, 1935, 1982, and 2094.

Main article: [Solar eclipse](#)

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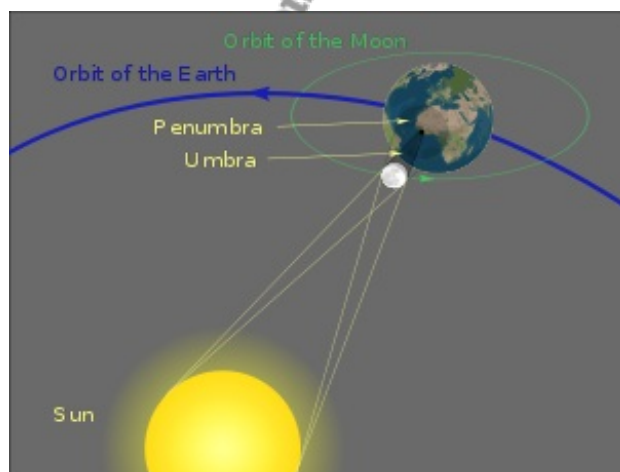
the viewer is inside the penumbra.



Each icon shows the view from the centre of its black spot, representing the moon (not to scale)

The [eclipse magnitude](#) is the fraction of the Sun's diameter that is covered by the Moon. For a total eclipse, this value is always greater than or equal to one. In both annular and total eclipses, the eclipse magnitude is the ratio of the angular sizes of the Moon to the Sun.<sup>[12]</sup>

Solar eclipses are relatively brief events that can only be viewed in totality along a relatively narrow track. Under the most favorable circumstances, a total solar eclipse can last for 7 minutes, 31 seconds, and can be viewed along a track that is up to 250 km wide. However, the region where a partial eclipse can be observed is much larger. The Moon's umbra will advance eastward at a rate of 1,700 km/h, until it no longer intersects the Earth's surface.



Geometry of a total solar eclipse (not to scale)

During a solar eclipse, the Moon can sometimes perfectly cover the Sun because its size is nearly the same as the Sun's when viewed from the Earth. A total solar eclipse is in fact an [occultation](#) while an annular solar eclipse is a [transit](#).

When observed at points in space other than from the Earth's surface, the Sun can be eclipsed by bodies other than the Moon. Two examples include when the crew of [Apollo 12](#) observed the [Earth to eclipse the Sun](#) in 1969 and when the [Cassini probe](#) observed [Saturn to eclipse the Sun](#) in 2006.

## Lunar eclipse

Main article: [Lunar eclipse](#)



The progression of a [lunar eclipse](#) from right to left. Totality is shown with the first two images. These required a longer [exposure time](#) to make the details visible.

Lunar eclipses occur when the Moon passes through the Earth's shadow. This happens only during a [full moon](#),

when the Moon is on the far side of the Earth from the Sun. Unlike a solar eclipse, an eclipse of the Moon can be observed from nearly an entire hemisphere. For this reason it is much more common to observe a lunar eclipse from a given location. A lunar eclipse lasts longer, taking several hours to complete, with totality itself usually averaging anywhere from about 30 minutes to over an hour.<sup>[13]</sup>

There are three types of lunar eclipses: penumbral, when the Moon crosses only the Earth's penumbra; partial, when the Moon crosses partially into the Earth's [umbra](#); and total, when the Moon crosses entirely into the Earth's umbra. Total lunar eclipses pass through all three phases. Even during a total lunar eclipse, however, the Moon is not completely dark. Sunlight refracted through the Earth's atmosphere enters the umbra and provides a faint illumination. Much as in a sunset, the atmosphere tends to more strongly scatter light with shorter wavelengths, so the illumination of the Moon by refracted light has a red hue,<sup>[14]</sup> thus the phrase 'Blood Moon' is often found in descriptions of such lunar events as far back as eclipses are recorded.<sup>[15]</sup>

## Historical record

Records of solar eclipses have been kept since ancient times. Eclipse dates can be used for [chronological dating](#) of historical records. A [Syrian](#) clay tablet, in the Ugaritic language, records a solar eclipse which occurred on March 5, 1223 B.C.,<sup>[16]</sup> while Paul Griffin argues that a stone in Ireland records an eclipse on November 30, 3340 B.C.<sup>[17]</sup> Positing classical-era astronomers' use of Babylonian eclipse records mostly from the 13th century BC provides a feasible and mathematically consistent<sup>[18]</sup> explanation for the Greek finding all three lunar mean motions (synodic, anomalistic, draconitic) to a precision of about one part in a million or better. Chinese historical records of solar eclipses date back over 4,000 years and have been used to measure changes in the Earth's rate of spin.<sup>[19]</sup>

By the 1600s, European astronomers were publishing books with diagrams explaining how lunar and solar eclipses occurred.<sup>[20][21]</sup> In order to disseminate this information to a broader audience and decrease fear of the consequences of eclipses, booksellers printed broadsides explaining the event either using the science or via astrology.<sup>[22]</sup>

## Some other planets and Pluto

### Gas giants



A picture of [Jupiter](#) and its moon [Io](#) taken by [Hubble](#). The black spot is Io's shadow.

The [gas giant](#) planets ([Jupiter](#),<sup>[23]</sup> [Saturn](#),<sup>[24]</sup> [Uranus](#),<sup>[25]</sup> and [Neptune](#))<sup>[26]</sup> have many moons and thus frequently display eclipses. The most striking involve [Jupiter](#), which has four large moons and a low [axial tilt](#), making eclipses more frequent as these bodies pass through the shadow of the larger planet. Transits occur with equal frequency. It is common to see the larger moons casting circular shadows upon Jupiter's cloudtops.

The eclipses of the [Galilean moons](#) by Jupiter became accurately predictable once their orbital elements were known. During the 1670s, it was discovered that these events were occurring about 17 minutes later than expected when Jupiter was on the far side of the Sun. [Ole Rømer](#) deduced that the delay was caused by the time needed for light to travel from Jupiter to the Earth. This was used to produce the first estimate of the [speed](#)



of light.<sup>[27]</sup>

On the other three [gas giants](#), eclipses only occur at certain periods during the planet's orbit, due to their higher [inclination](#) between the orbits of the moon and the [orbital plane](#) of the planet. The moon [Titan](#), for example, has an orbital plane tilted about 1.6° to Saturn's equatorial plane. But Saturn has an axial tilt of nearly 27°. The orbital plane of Titan only crosses the line of sight to the Sun at two points along Saturn's orbit. As the orbital period of Saturn is 29.7 years, an eclipse is only possible about every 15 years.

The timing of the Jovian satellite eclipses was also used to calculate an observer's [longitude](#) upon the Earth. By knowing the expected time when an eclipse would be observed at a standard longitude (such as [Greenwich](#)), the time difference could be computed by accurately observing the local time of the eclipse. The time difference gives the longitude of the observer because every hour of difference corresponded to 15° around the Earth's equator. This technique was used, for example, by [Giovanni D. Cassini](#) in 1679 to re-map [France](#).<sup>[28]</sup>

## Mars

Main article: [Transit of Phobos from Mars](#)

On [Mars](#), only partial solar eclipses ([transits](#)) are possible, because neither of its moons is large enough, at their respective orbital radii, to cover the Sun's disc as seen from the surface of the planet. Eclipses of the moons by Mars are not only possible, but commonplace, with hundreds occurring each Earth year. There are also rare occasions when Deimos is eclipsed by Phobos.<sup>[29]</sup> Martian eclipses have been photographed from both the surface of Mars and from orbit.

## Pluto

Main article: [Solar eclipses on Pluto](#)

[Pluto](#), with its proportionately largest moon [Charon](#), is also the site of many eclipses. A series of such mutual eclipses occurred between 1985 and 1990.<sup>[30]</sup> These daily events led to the first accurate measurements of the physical parameters of both objects.<sup>[31]</sup>

## Mercury and Venus

Eclipses are impossible on [Mercury](#) and [Venus](#), which have no moons. However, both have been observed to transit across the face of the Sun. There are on average 13 transits of Mercury each century. Transits of Venus occur in pairs separated by an interval of eight years, but each pair of events happen less than once a century.<sup>[32]</sup>

## Eclipsing binaries

A [binary star](#) system consists of two stars that orbit around their common [centre of mass](#). The movements of both stars lie on a common [orbital plane](#) in space. When this plane is very closely aligned with the location of an observer, the stars can be seen to pass in front of each other. The result is a type of extrinsic [variable star](#) system called an [eclipsing binary](#).

The maximum [luminosity](#) of an eclipsing binary system is equal to the sum of the luminosity contributions from the individual stars. When one star passes in front of the other, the luminosity of the system is seen to decrease. The luminosity returns to normal once the two stars are no longer in alignment.<sup>[33]</sup>

The first eclipsing binary star system to be discovered was [Algol](#), a star system in the constellation [Perseus](#). Normally this star system has a [visual magnitude](#) of 2.1. However, every 2.867 days the magnitude decreases to 3.4 for more than nine hours. This is caused by the passage of the dimmer member of the pair in front of the brighter star.<sup>[34]</sup> The concept that an eclipsing body caused these luminosity variations was introduced by [John](#)

Goodricke in 1783.<sup>[35]</sup>

## See also

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### Lunar eclipses

Lists of lunar eclipses

Lunar eclipses

by century

Lunar eclipses

by [Saros series](#)





Partial eclipses  
1950–2049



Total eclipses  
1950–2049  
Penumbral eclipses

See also

- [Danjon scale](#)
- [Gamma](#)

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[Lists of eclipses](#)

By century  
[Saros series](#)  
Visibility  
Historical



Total/hybrid eclipses

→ symbol denotes  
next total/hybrid eclipse



Annular eclipses

→ symbol denotes  
next annular eclipse



Partial eclipses

→ symbol denotes  
next partial eclipse

Other bodies

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