

# The Impact of Accretion Disk Winds on the Optical Spectra of Cataclysmic Variables

J. H. Matthews, C. Knigge, K. S. Long, S. A. Sim & N. Higginbottom

13 January 2015

## ABSTRACT

Many high-state non-magnetic cataclysmic variables (CVs) exhibit blue-shifted absorption or P-Cygni profiles in ultraviolet (UV) resonance lines, such as C iv 1550Å. These features imply the existence of powerful accretion disk winds in CVs. Here, we use an improved version of our Monte Carlo ionization and radiative transfer code to investigate whether disk wind models that produce realistic UV line profiles are also likely to generate observationally significant recombination line and continuum emission in the *optical* waveband. In particular, we test whether outflows may be responsible for the single-peaked emission line profiles often seen in high-state CVs and for the weakness of the Balmer absorption edge (relative to simple models of optically thick accretion disks). We find that a standard disk wind model that is successful in reproducing the UV spectra of CVs also leaves a noticeable imprint on the optical spectrum, particularly for systems viewed at high inclination. The strongest optical wind-formed recombinations lines are H $\alpha$  and He II  $\lambda$ 4686Å although the outflow also produces emission in the Ly- $\alpha$  and He II  $\lambda$ 1640Å UV lines. However, all of the optical lines produced by this default wind model are double-peaked, and the effect of the wind on the Balmer jump is modest. We demonstrate that a higher-density outflow model displays reasonable verisimilitude with the optical spectrum of RW Trianguli, producing all the observed lines. In addition, this model produces recombination continuum sufficient to fill in the Balmer jump, although it also predicts stronger-than-observed collisionally excited UV emission lines. Our revised model also significantly alters the spectral shape of the continuum and displays a strong inclination dependence in its emission properties. No single-peaked emission is seen, although we observe a narrowing of the double-peaked emission lines from the base of the wind. We suggest that a non-homogenous, ‘clumpy’ wind may provide a possible solution. On the basis of these results, we suggest that winds can modify, and perhaps even dominate, the line and continuum emission from accreting systems.

## 1 INTRODUCTION

Cataclysmic variables (CVs) are systems in which a white dwarf accretes matter from a donor star via Roche-lobe overflow. In non-magnetic systems this accretion is mediated by a Keplerian disk around the white dwarf (WD). Nova-like variables (NLs) are a subclass of CVs in which the accretion disk is always in a relatively high accretion rate state ( $\dot{M} \sim 10^{-8} M_{\odot} \text{ yr}^{-1}$ ). This makes NLs an excellent laboratory for studying the properties of steady-state accretion disks.

It has been known for a long time that winds emanating from the accretion disk are important in shaping the ultraviolet (UV) spectra of high-state CVs (Heap et al. 1978; Greenstein & Oke 1982). The most spectacular evidence for such outflows are the P-Cygni-like profiles seen in UV resonance lines such as C iv  $\lambda$ 1550Å (see e.g. Cordova & Mason 1982). Considerable effort has been spent over the years on understanding and modelling these UV features (e.g. Drew & Verbunt 1985; Mauche & Raymond 1987; Drew 1987; Shlosman & Vitello 1993; [hereafter SV93]; Knigge,

Woods & Drew 1995 [hereafter KWD95]; Knigge & Drew 1997; Knigge et al. 1997; Long & Knigge 2002 [hereafter LK02], Noebauer et al. 2010; Puebla et al. 2011). The basic picture emerging from these efforts is of a slowly accelerating, moderately collimated bipolar outflow that carries away  $\simeq 1\% - 10\%$  of the accreting material. State-of-art-simulations of line formation in this type of disk wind can produce UV line profiles that are remarkably similar to observations.

Much less is known about the effect of these outflows on the optical spectra of high-state CVs. These spectra are typically characterized by H and He emissions superposed on a blue continuum. In many cases, and particularly in the SW Sex subclass of NLs (Honeycutt et al. 1986; Dhillon & Rutten 1995), these lines are single-peaked. This is contrary to theoretical expectations for lines formed in accretion disks, which are predicted to be double-peaked (Smak 1981; Horne & Marsh 1986). *Low-state* CVs (dwarf novae in quiescence) can, in fact, exhibit such double-peaked lines (Marsh & Horne 1990).

Murray & Chiang (1996, 1997; hereafter referred to collectively as MC96) have shown that the presence of disk winds may offer a natural explanation for the single-peaked optical emission lines in high-state CVs, since they can strongly affect the radiative transfer of line photons. Strong support for a significant wind contribution to the optical emission lines comes from observations of eclipsing systems. There, the single-peaked lines are often only weakly eclipsed, and a significant fraction of the line flux remains visible even near mid-eclipse (REF). This points to line formation in a spatially extended region, such as a disk wind (see figure 1). Further evidence for a wind contribution to the optical lines comes from isolated observations of P-Cygni-like line profiles even in optical lines, such as H $\alpha$  and He I  $\lambda$ 5876 (Patterson et al. 1996; Ringwald & Naylor 1998; Kafka & Honeycutt 2004).

Could disk winds also have an impact on the UV/optical *continuum* of high-state CVs? This continuum is usually thought to be dominated by the accretion disk and modelled by splitting the disk into a set of concentric, optically thick, non-interacting annuli following the standard  $T_{eff}(R) \propto R^{-3/4}$  radial temperature distribution (Shakura & Sunyaev 1973). In such models, each annulus is taken to emit either as a blackbody or, perhaps more realistically, as a stellar/disk atmosphere model (REF). In the latter case, the local surface gravity,  $\log g(R)$ , is assumed to be set solely by the accreting WD, since self-gravity is negligible in CV disks.

Attempts to fit the observed spectral energy distributions (SEDs) of high-state CVs with such models have met with mixed success. In particular, the SEDs predicted by most stellar/disk atmosphere models are too blue in the UV (Wade 1988; Long et al. 1991, 1994) (Other refs here?) and exhibit stronger-than-observed Balmer jumps in absorption (Wade 1984; Haug 1987; La Dous 1989a; Knigge et al. 1998a). One possible explanation for these problems is that these models fail to capture all of the relevant physics. Indeed, it has been argued that a self-consistent treatment can produce better agreement with observational data (e.g. Shaviv et al. 1991; but also more recent work by Idan et al. 2010). However, an alternative explanation, suggested by Knigge et al. (1998b), is that recombination continuum emission from the base of the disk wind might fill in the disk’s Balmer absorption edge and flatten the UV spectrum.

Here, we carry out Monte Carlo radiative transfer simulations in order to assess the likely impact of accretion disk winds on the optical spectra of high-state CVs. More specifically, our goal is to test whether disk winds of the type developed to account for the UV resonance lines would also naturally produce significant amounts of optical line and/or continuum emission. In order to achieve this, we have implemented the ‘macro-atom’ approach developed by Lucy (2002, 2003) into the Monte Carlo ionization and radiative transfer code described by LK02 (a process initiated by SDL05). With this upgrade, the code is able to deal correctly with processes involving excited levels, such as the recombination emission produced by CV winds.

The remainder of this paper is organized as follows. In Section 2 we briefly describe the code and the newly implemented macro-atom approach. In Section 3, we describe the kinematics and geometry of our disk wind model. In Section 4, we present spectra simulated from the bench-

mark model employed by LK02, and in Section 5, we carry out a limited sensitivity study to explore how the predicted spectra change as a function of some key parameters in our model. In Section 6, we summarize our findings.

## 2 PYTHON: A MONTE CARLO IONIZATION AND RADIATIVE TRANSFER CODE

PYTHON is a Monte Carlo ionization and radiative transfer code which uses the Sobolev approximation to treat line transfer. It has already been described extensively in LK02, SDL05 and H13, so here we provide only a brief summary of its operation, focusing particularly on new aspects of our implementation of macro-atoms into the code.

### 2.1 Basics

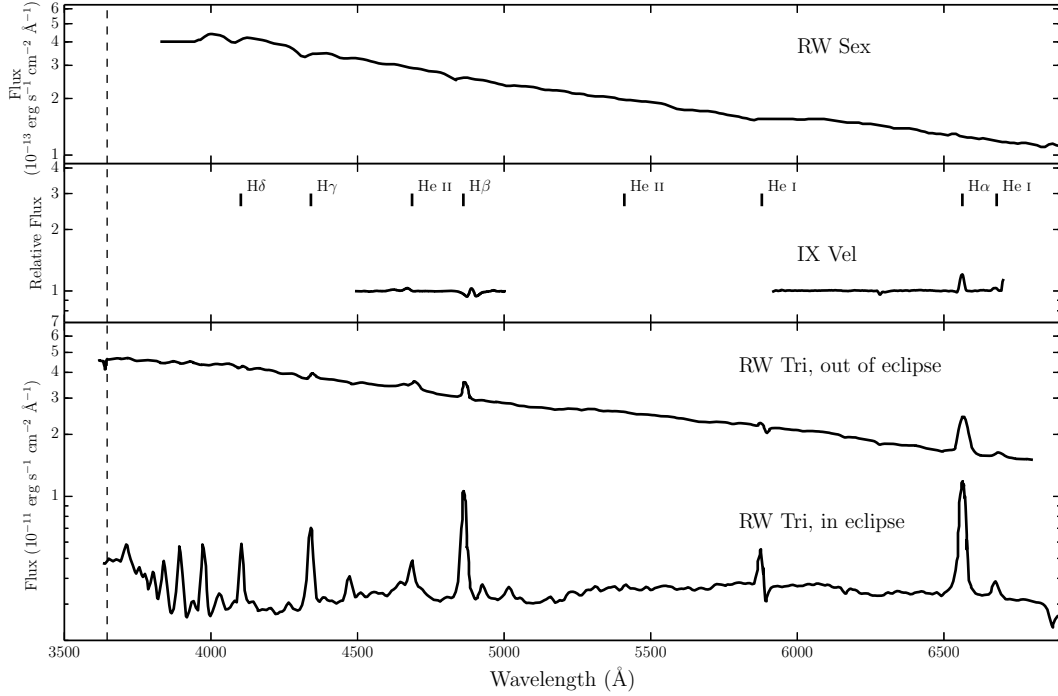
PYTHON operates in two distinct stages. First, the ionization state, level populations and temperature structure are calculated. This is done iteratively, by propagating several populations of Monte Carlo energy quanta (‘photons’) through a model wind. The geometric and kinematic properties of the outflow are specified on a pre-defined spatial grid. In each of these iterations (‘ionization cycles’), the code records estimators that characterize the radiation field in each grid cell. At the end of each ionization cycle, a new electron temperature is calculated that more closely balances heating and cooling in the plasma. The radiative estimators and updated electron temperature are then used to revise the ionization state of the wind, and a new ionization cycle is started. The process is repeated until heating and cooling are balanced throughout the wind.

This converged model is used as the basis for the second set of iterations (‘spectral cycles’). In these, the emergent spectrum over the desired spectral range is synthesized by tracking populations of energy packets through the wind and computing the emergent spectra at a number of user-specified viewing angles.

PYTHON is designed to operate in a number of different regimes, both in terms of the scale of the system, and in terms of the characteristics of the underlying radiation field. It was originally developed by LK02 in order to model the UV spectra of CVs with a simple biconical disk wind model. Sim et al. (2005; hereafter SDL05) used the code to model Brackett and Pfund line profiles of H in young-stellar objects (YSOs). As part of this effort, they implemented a ‘macro-atom’ mode (see below) in order to correctly treat hydrogen recombination lines with PYTHON. Finally, Higginbottom et al. (2013; hereafter H13) used PYTHON to model broad absorption line (BAL) QSOs. For this application, an improved treatment of ionization was implemented, so that the code is now capable of dealing with arbitrary photo-ionizing SEDs, including non-thermal and multi-component ones.

### 2.2 Ionization and Excitation: ‘Simple Atoms’

Prior to SDL05, the relative ionization fractions for all atomic species were estimated via the modified Saha equation (Mazzali & Lucy 1993)



**Figure 1.** Optical spectra of three nova-like variables: RW Sex (top; Beuermann et al. 1992), IX Vel (top middle; Beuermann & Thomas 1990) and RW Tri in and out of eclipse (bottom two panels; Groot et al. 2004). The data have been digitized from the respective publications and a few of the more prominent emission and absorption lines are marked. These systems have inclinations of  $30^\circ$ ,  $60^\circ$  and  $70^\circ$  respectively. The trend of increasing Balmer line emission with inclination can be seen. In RW Tri strong single-peaked emission in the Balmer lines is seen even in eclipse, indicating that the lines may be formed in a spatially extensive disk wind, and there is even a suggestion of a (potentially wind-formed) recombination continuum in the eclipsed spectrum. We have attempted to show each spectrum over a similar dynamic range.

$$\frac{n_{j+1}n_e}{n_j} = W[\xi + W(1 - \xi)] \left( \frac{T_e}{T_R} \right)^{1/2} \left( \frac{n_{j+1}n_e}{n_j} \right)^*_{T_R}, \quad (1)$$

Here, the ‘starred’ term on the right represents abundances computed with the Saha equation at temperature  $T_R$ , but using actual partition functions from the dilute blackbody approximation.  $W$  is an effective dilution factor,  $\xi$  is the fraction of recombinations going directly to the ground state, and  $T_R$  and  $T_e$  are the radiation and electron temperature respectively. This simple ionization scheme produces reasonable results when the photoionizing SED can be approximated by a dilute blackbody. This is the case for high-state CVs. (As noted above, an improved, but more complex treatment of ionization that is appropriate for more complex SEDs is described in H13.)

Similarly, the relative excitation fractions within each ionization stage of a given species were estimated via a modified Boltzmann equation,

$$\frac{n_{jk}}{n_j} = \frac{Wg_k}{z_j(T_R)} \exp(-E_k/kT_R), \quad (2)$$

where  $n_{jk}$  is the population of level  $k$  in ionic stage  $j$ ,  $E_k$  is the energy difference between level  $k$  and the ground state, and  $z_j(T_R)$  is the partition function of ionic stage  $j$ .

Finally, PYTHON modelled all bound-bound processes as transitions within a simple two-level atom (see e.g. Mihalas 1982). This approach works reasonable well for resonance lines, such as C IV  $\lambda 1550\text{\AA}$ , in which the lower level is the ground state. However, it is a poor approximation for

many other transitions, particularly those where the upper level is primarily populated from above. Thus an improved method for estimating excited level populations and simulating line transfer is needed in order to model recombination lines and continua.

### 2.3 Ionization and Excitation: Macro-Atoms

Lucy (2002, 2003; hereafter L02, L03) has shown that it is possible to calculate the emissivity of a gas in statistical equilibrium accurately by quantising matter into ‘macro-atoms’, and radiant and kinetic energy into indivisible energy packets (r- and k- packets, respectively). His macro-atom scheme allows for all possible transition paths from a given level and provides a full non-local thermodynamic equilibrium (NLTE) solution for the level populations based on Monte Carlo estimators. The macro-atom technique has already been used to model Wolf-Rayet star winds (Sim 2004), AGN disk winds (Sim et al. 2008; Tatum et al. 2012), supernovae (Kasen et al. 2006; Kerzendorf & Sim 2014) and YSOs (SDL05). A full description of the can be found in L02 and L03.

Briefly, macro-atom NLTE level populations and ionization fractions are calculated by solving the statistical equilibrium equations between each pair of levels. For example, the bound-bound de-excitation rate,  $\mathcal{R}_{ul}$ , in an ion is given by

$$\mathcal{R}_{ul} = \beta_{lu}A_{ul}n_u + \beta_{lu}B_{ul}n_uJ_{lu} + C_{ul}n_un_e, \quad (3)$$

where  $u$  and  $l$  denote the upper and lower levels,  $C$  represents the collisional rate coefficients and  $A$  and  $B$  represent the Einstein coefficients. The quantity  $\beta_{lu}$  is the probability that a given line photon will escape the Sobolev region. Finally,  $J$  is the Monte Carlo estimator for the mean intensity in the line, as recorded during the photon propagation. The corresponding excitation rate is then

$$\mathcal{R}_{lu} = \beta_{lu} B_{lu} n_l J_{lu} + C_{lu} n_l n_e, \quad (4)$$

where we implicitly include a correction for stimulated emission in our definition of the estimate for  $J_{ul}$ .

In our implementation of the macro-atom approach, we also explicitly take into account the photoionization and collisional ionization rates between a lower level,  $l$ , and the continuum (or, in the case of ions with more than one bound electron, the ground state of the upper ion),  $\kappa$ ,

$$\mathcal{R}_{l\kappa} = n_l \int_{\nu_0}^{\infty} \frac{4\pi J_{\nu} \sigma_{\nu}}{h\nu} d\nu + C_{l\kappa} n_l n_e. \quad (5)$$

Here,  $\sigma_{\nu}$  is the photoionization cross section, and  $J_{\nu}$  is the mean intensity. The corresponding recombination rate is given by

$$\mathcal{R}_{\kappa l} = \alpha_{\kappa l} n_{\kappa} n_e + C_{\kappa l} n_{\kappa} n_e, \quad (6)$$

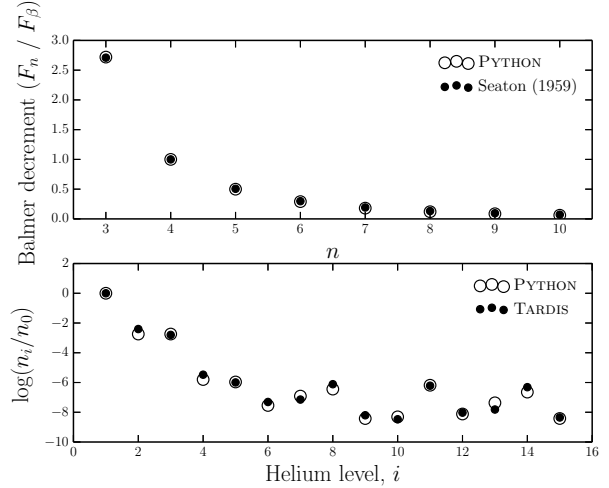
where  $\alpha_{\kappa l}$  is the radiative recombination coefficient to level  $l$ . This treatment means that radiative and collisional rates to and from all levels are considered when calculating both the ionization state and the level populations, although we neglect ionization directly to excited levels of the upper ion. The van Regemorter (1962) approximation is used for collisional transitions, and thus we assume that collisions between radiatively forbidden transitions are not important in determining level populations. This is a reasonable approximation in the regimes considered here.

## 2.4 Ionization and Excitation: A Hybrid Approach

SDL05 implemented a macro-atom treatment of Hydrogen in PYTHON and used this to predict the observable properties of a pure Hydrogen wind model for YSOs. Our goal here is to simultaneously model the optical and ultraviolet spectra of high-state CVs. Since the optical spectra are dominated by H and He recombination lines, both of these species need to be treated as macro-atoms. The UV spectra, on the other hand, are dominated by resonance lines associated with metals. This means we need to include these species in our models, but they can be treated with our (much faster) simple-atom approach. We have therefore implemented a hybrid ionization and excitation scheme into PYTHON. Any atomic species can now be treated either in our simple-atom approximation or with the full macro-atom machinery. For our CV models, we treat H and He as macro-atoms and all metals as simple-atoms.

## 2.5 Atomic Data

We generally use the same atomic data as H13, which is an updated version of that described in LK02. In addition, we follow SDL05 in treating Hydrogen as a 20-level atom, where each level is defined by the principal quantum number,  $n$ .



**Figure 2.** *Top Panel:* “Case B” Balmer decrements computed with PYTHON compared to analytic calculations by Seaton (1959). Both calculations are calculated at  $T_e = 10,000\text{K}$ . (see Osterbrock 1989 for a discussion of this commonly used approximation). *Bottom Panel:* a comparison of Helium I level populations (the most complex ion we currently treat as a macro-atom) between PYTHON and TARDIS models. The calculation is conducted with physical parameters of  $n_e = 5.96 \times 10^4 \text{ cm}^{-3}$ ,  $T_e = 30,600\text{K}$ ,  $T_R = 43,482\text{K}$  and  $W = 9.65 \times 10^{-5}$ . Considering the two codes use different atomic data and TARDIS, unlike PYTHON, currently has a complete treatment of collisions between radiatively forbidden transitions, the factor of  $< 2$  agreement is encouraging.

For the macro-atom treatment of Helium, we have added additional level and line information required from TOPBASE (Badnell et al. 2005). He II is treated in much the same way as H, but with 10 levels. He I has larger energy differentials between different l-subshells and triplet and singlet states. Thus, we still include levels up to  $n = 10$  but explicitly treat the  $l$  and  $s$  sub-orbitals as distinct levels instead of assuming they are perfectly ‘l-mixed’. This allows us to model the singlet and triplet He I lines that are ubiquitous in the optical spectra of CVs (e.g. Dhillon 1996).

## 2.6 Code Validation and Testing

PYTHON has been tested against a number of radiative transfer and photoionization codes. LK02 and H13 conducted comparisons of ionization balance with CLOUDY (Ferland et al. 2013), demonstrating excellent agreement. We have also carried out comparisons of ionization and spectral synthesis with the supernova code TARDIS. TARDIS is described by Kerzendorf & Sim (2014), and the spectral comparisons can be found therein. For the effort reported here, we have additionally carried out tests of the macro-atom scheme in PYTHON. Figure 2 shows two of these tests. In the top panel, we compare the Balmer series emissivities as predicted by PYTHON in the l-mixed Case B limit against the analytical calculations by Seaton (1959). In the bottom panel, we compare PYTHON and TARDIS predictions of He I level populations for a particular test case. Agreement is excellent for both H and He.



### 3 DESCRIBING THE SYSTEM AND ITS OUTFLOW

PYTHON includes several different kinematic models of accretion disk winds, as well as different options for describing the physical and radiative properties of the wind-driving system under consideration. Most of these features have already been discussed in LK02 and H13, so below we only briefly recount the key aspects of the particular system and wind model used in the present study.

#### 3.1 Wind Geometry and Kinematics

We adopt the kinematic disk wind model developed by Schlosman & Vitello (1993; hereafter SV93). A schematic of this model is shown in figure 3. In this parametrization, a smooth, biconical disk wind emanates from the accretion disk between radii  $r_{min}$  and  $r_{max}$ . The covering fraction of the outflow is also controlled by the inner and outer opening angles of the wind,  $\theta_{min}$  and  $\theta_{max}$ , and the launch angle of the other streamlines is given by

$$\theta(r_0) = \theta_{min} + (\theta_{max} - \theta_{min}) \left( \frac{r_0 - r_{min}}{r_{max} - r_{min}} \right)^\gamma, \quad (7)$$

where  $r_0$  is the launch radius of the streamline.

The poloidal (non-rotational) velocity field of the wind,  $v_l$ , is given by

$$v_l = v_0 + [v_\infty(r_0) - v_0] \frac{(l/R_v)^\alpha}{(l/R_v)^\alpha + 1}, \quad (8)$$

where  $l$  is the poloidal distance along a particular wind streamline. The terminal velocity along a streamline,  $v_\infty$ , is set to a fixed multiple of the escape velocity at the launch point. The launch velocity from the disk surface,  $v_0$ , is assumed to be constant (set to  $6 \text{ km s}^{-1}$ ). Once the wind is launched, it accelerates, reaching half of its terminal velocity at  $l = R_v$ . The velocity law exponent  $\alpha$  controls how quickly the wind accelerates. Larger values of  $\alpha$  cause the main region of acceleration to occur close to  $R_v$ , whereas smaller values correspond to fast acceleration close to the disk (see figure 4).

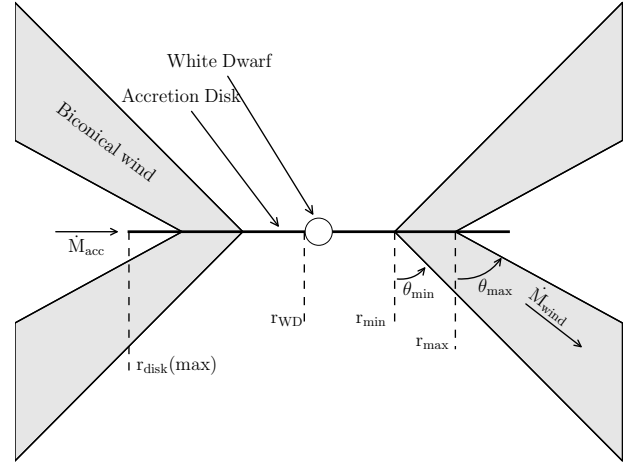
The density at position  $(r, z)$  in the wind,  $\rho(r, z)$ , is calculated from the mass continuity equation, yielding

$$\rho(r, z) = \frac{r}{r_0} \frac{dr}{dr_0} \frac{\phi(r_0)}{v_z(r, z)}. \quad (9)$$

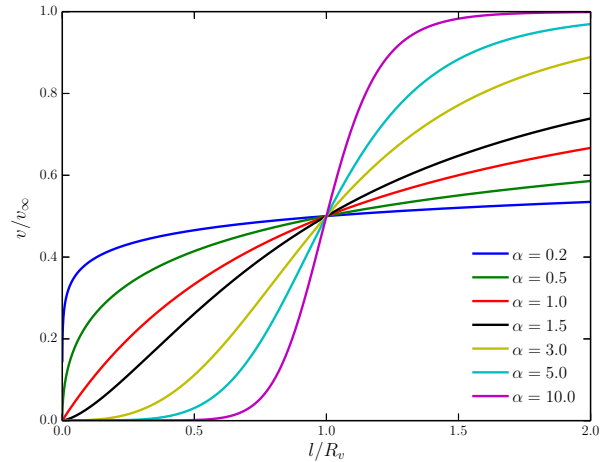
Here,  $\phi(r_0)$  is the mass-loss rate per unit area at  $r_0$ , and  $v_z$  is the vertical velocity component. We take  $\phi(r_0)$  to be uniform across the wind-launching area of the disk and normalize it by matching its integral to the user-specified total mass-loss rate,  $\dot{M}_{wind}$ . **CK: I'm still not sure I understand the formula SV93 use to define  $\phi$ , and hence that this description is actually strictly correct.... Not sure what to do about that though...**

#### 3.2 Sources and Sinks of Radiation

The net photon sources in our CV model are the accretion disk, the WD and, in principle, a boundary layer with user-defined temperature and luminosity. All of these radiating bodies are taken to be optically thick, and photons striking them are assumed to be destroyed instantaneously. The



**Figure 3.** Cartoon illustrating the geometry and kinematics of the benchmark CV wind model.



**Figure 4.** The adopted poloidal velocity law for various values of the acceleration exponent,  $\alpha$ . **JM: KSL has suggested we ditch this**

secondary star is not included as a radiation source, but is included as an occulting body. This allows us to model eclipses. Finally, emission from the wind itself is also accounted for, but note that we assume the outflow to be in radiative equilibrium. Thus all of the heating of the wind, as well as its emission, is ultimately powered by the radiation field of the net photon sources in the simulation. In the following sections, we will describe our treatment of these system components in slightly more detail.

##### 3.2.1 Accretion Disk

PYTHON has some flexibility when treating the accretion disk as a source of photons. The disk is broken down into annuli such that each annulus contributes an equal amount to the bolometric luminosity. We take the disk to be geometrically thin, but optically thick and thus adopt the temperature profile of a standard Shakura & Sunyaev (1973)  $\alpha$ -disk. An annulus can then be treated either as a black-

body with the corresponding effective temperature or as a stellar atmosphere model with the appropriate surface gravity and effective temperature. Here, we use blackbodies during the ionization cycles and to compute our Monte Carlo estimators. However, during the spectral synthesis stage of the simulation we use stellar atmosphere models. This produces more realistic model spectra and allows us to test if recombination emission from the wind base can fill in the Balmer jump, which is always in absorption in these models. Our synthetic stellar atmosphere spectra are calculated with SYNSPEC<sup>1</sup> from either Kurucz (Kurucz 1991) atmospheres (for  $T_{eff} \leq 50,000$  K) or from TLUSTY models (Hubeny & Lanz 1995) (for  $T_{eff} > 50,000$  K).

### 3.2.2 White Dwarf

The WD at the center of the disk is always present as a spherical occulting body with radius  $R_{WD}$  in PYTHON CV models, but it can also be included as a source of radiation. In the models presented here, we treat the WD as a blackbody radiator with temperature  $T_{WD}$  and luminosity  $L_{WD} = 4\pi R_{WD}^2 \sigma T_{WD}^4$ .

### 3.2.3 Boundary Layer

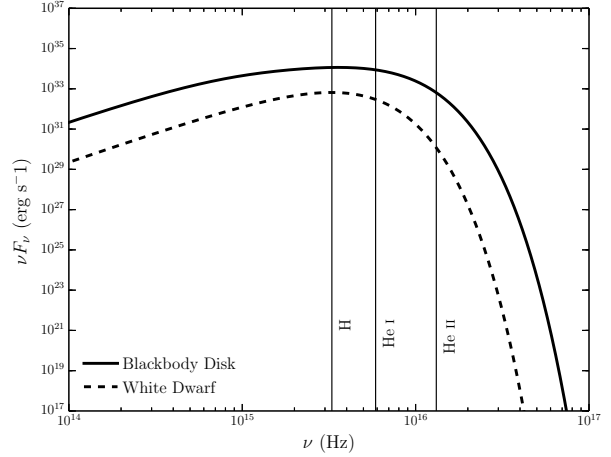
It is possible to include radiation from a boundary layer (BL) between the disk and the WD. In PYTHON, the BL is described as a blackbody with a user-specified effective temperature and luminosity. In the models presented here, we have followed LK02 in setting the BL luminosity to zero. However, we have confirmed that the addition of a BL with  $L_{BL} = 0.5L_{acc}$  and temperatures in the range  $80\text{kK} \leq T_{BL} \leq 200\text{kK}$  would not change any of our main conclusions.

### 3.2.4 Secondary Star

The donor star is included in the system as a pure radiation sink, i.e. it does not emit photons, but absorbs any photons that strike its surface. The secondary is assumed to be Roche-lobe filling, so its shape and size are defined by setting the mass ratio of the system,  $q = M_2/M_{WD}$ . The inclusion of the donor star as an occulting body allows us to model eclipses of the disk and the wind. For this purpose, we define orbital phase such that  $\Phi_{orb} = 0$  as the inferior conjunction of the secondary (i.e. mid-eclipse for  $i \simeq 90^\circ$ ).

## 4 A BENCHMARK DISK WIND MODEL

Our main goal is to test whether the type of disk wind model that has been successful in explaining the UV spectra of CVs could also have significant impact on the optical continuum and emission line spectra of these systems. In order to set a benchmark, we therefore begin by investigating one of the fiducial CV wind models that was used by SV93 and LK02 to simulate the UV spectrum of a typical high-state system. The specific parameters for this model (Model A) are listed in Table 1. A key point is that the wind mass-loss rate in



**Figure 5.** The spectral energy distribution (SED) of the White Dwarf (dashed line) and blackbody accretion disk used in the ionization calculation (black solid line). The positions of the Hydrogen and Helium ionization edges are marked with vertical lines. **JM: KSL has suggested we ditch this**

Model Parameters

Parameter	Model A	Model B
$M_{WD}$	$0.8M_\odot$	
$R_{WD}$	$7 \times 10^8 \text{ cm}$	
$T_{WD}$	$40,000 \text{ K}$	
$M_2$	$0.6M_\odot$	
$\dot{M}_{acc}$	$10^{-8} M_\odot \text{ yr}^{-1}$	
$\dot{M}_{wind}$	$10^{-9} M_\odot \text{ yr}^{-1}$	
$r_{min}$	$4R_{WD}$	
$r_{max}$	$12R_{WD}$	
$\theta_{min}$	$20.0^\circ$	
$\theta_{max}$	$65.0^\circ$	
$\gamma$	1	
$v_\infty$	$3v_{esc}$	
$R_v$	$100R_{WD}$	$142.9R_{WD}$
$\alpha$	1.5	4

**Table 1.** Parameters used for the geometry and kinematics of the benchmark CV model (Model A), which is optimized for the UV band, and a model which is optimized for the optical band and described in section 6 (Model B). For Model B, only parameters which are altered are given - otherwise the Model A parameter is used. **JM: this needs some stuff adding, as is mentioned in certain places**

this model is set to 10% of the accretion rate through the disk.

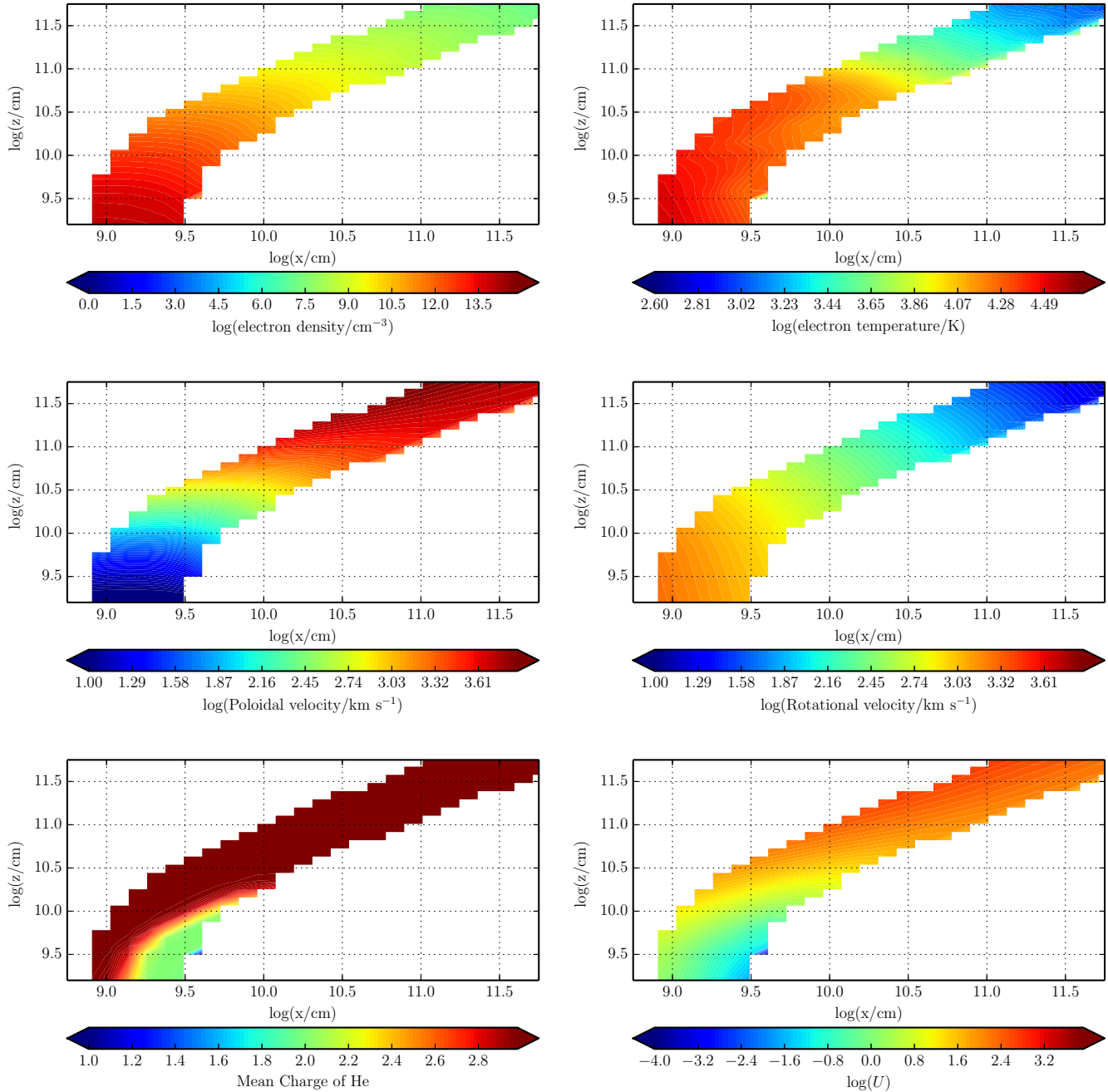
### 4.1 Physical Structure and Ionization State

Figure 6 shows the physical and ionization structure of the benchmark disk wind model. The ionization parameter shown in the bottom right panel is given by

$$U = \frac{Q_H}{4\pi R^2 n_{HC}} \quad (10)$$

$$= \frac{4\pi}{n_{HC}} \int_{13.6\text{eV}}^{\infty} \frac{J_\nu d\nu}{h\nu} \quad (11)$$

<sup>1</sup> <http://nova.astro.umd.edu/Synspec43/synspec.html>



**Figure 6.** The physical properties of the wind- note the logarithmic scale. Near the disk plane the wind is dense, with low poloidal velocities. As the wind accelerates it becomes less dense and more highly ionized. The dominant Helium ion is almost always He III, apart from in a small portion of the wind at the base, which is partially shielded from the inner disk.

where  $Q_H$  is the number of hydrogen-ionizing photons emitted by the illuminating source per second,  $R$  is the distance to the source,  $n_H$  is the local hydrogen number density and  $c$  is the speed of light.  $J_\nu$  is the Monte Carlo estimator for the mean intensity in the cell. The ionization parameter is a useful measure of the ionization state of a plasma, as it evaluates the ratio of the integrated ionizing radiation field and the local Hydrogen density.

There is an obvious drop-off in density and temperature with distance away from the disk, so any line formation process that scales as  $\rho^2$  - i.e. recombination and collisionally

excited emission - should be expected to operate primarily in the dense base of the outflow. Moreover, a comparison of the rotational and poloidal velocity fields shows that rotation dominates in the near-disk regime, while outflow dominates further out in the wind.

The ionization equation used in the “simple atom” approach used in LK02 (see section 2.2) should be a reasonable approximation to the photoionization equilibrium in the benchmark wind model. Even though the macro-atom treatment of H and He does affect the computation of the overall ionization equilibrium, we would expect the resulting ion-

ization state of the wind to be quite similar to that found in LK02. The bottom panels in figure 6 confirm that this is the case. In particular, Helium is fully ionized throughout most of the outflow, except for a small region near the base of the wind, which is shielded from the photons produced by the hot inner disk. In line with the results shown in LK02, we also find that C IV is the dominant Carbon ion throughout the wind, resulting in a substantial absorbing column across a large range of velocities. As we shall see, this produces the broad, deep and blue-shifted C IV  $\lambda 1550\text{\AA}$  absorption line that is usually the most prominent wind-formed feature in the UV spectra of low-inclination nova-like CVs.

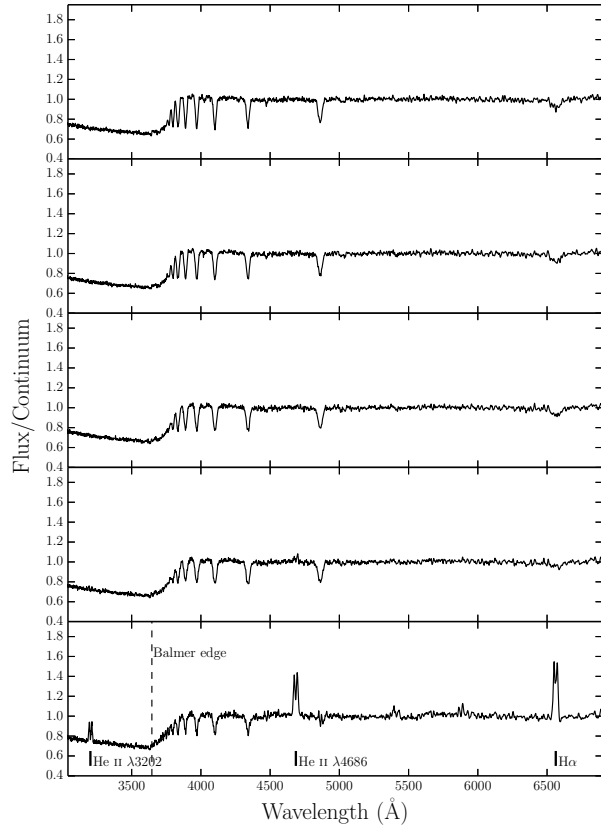
#### 4.2 Synthetic Spectra

We begin by verifying that the benchmark model still produces UV spectra that resemble those observed in CVs. We do expect this to be the case, since the ionization state of the wind has not changed significantly from that computed in LK02 (see section 4.1). The left column of panels in figure 7 shows that this expectation is met: all of the strong metal resonance lines - notably N V  $\lambda 1240\text{\AA}$ , Si IV  $\lambda 1400\text{\AA}$  and C IV  $\lambda 1550\text{\AA}$  - are present and exhibit clear P-Cygni profiles at intermediate inclinations. In addition, however, we now also find that the wind produces significant Lyman- $\alpha$  and He II  $\lambda 1640\text{\AA}$  emission lines.

Figure 7 shows the corresponding optical spectra produced for the benchmark model, and these do exhibit some emission lines associated with H and He. We see a general trend from absorption lines to emission lines with increasing inclination, as one might expect from our wind geometry. This trend is consistent with observations, as can be seen in Figure 1. However, it is clear that this particular model does not produce all the lines seen in observations of high-state CVs. The higher order Balmer series lines too weak to overcome the intrinsic absorption from the disk atmosphere, and the wind fails to produce any emission at intermediate inclinations, which contrasts with (for example) the observed behaviour of IX Vel (Beuermann & Thomas 1990).

Our models do not exhibit P-Cygni profiles in the optical lines. This is perhaps not surprising. LK02 and SV93 originally designed such models to reproduce the UV line profiles. Thus, most of the wind has an ionization parameter of  $\log U \sim 2$  (see figure 6). This means Hydrogen and Helium are fully ionized throughout much of the wind and is successful in producing recombination features. However, it leads to absorbing columns in H $\alpha$  of  $N_H \sim ??$  - far too low to produce noticeable blue shifted absorption. While we suspect that the winds that exhibit such profiles must lie in a different ionization regime, requiring significantly different disk SEDs or mass-loss rates to a ‘standard’ high-state CV. However, we note it is difficult for this idea to exist within a line-driven wind framework.

The emissivity of these recombination features scales as  $\rho^2$ , meaning that they form almost entirely in the dense base of the wind, just above the accretion disk. Here, the velocity field of the wind is still dominated by rotation, rather than outflow, which accounts for the double-peaked shape of the lines. In principle, lines formed in this region can still be single peaked, since the existence of a poloidal velocity *gradient* changes the local escape probabilities (MC96). However, as discussed further in section ??, the radial velocity shear in

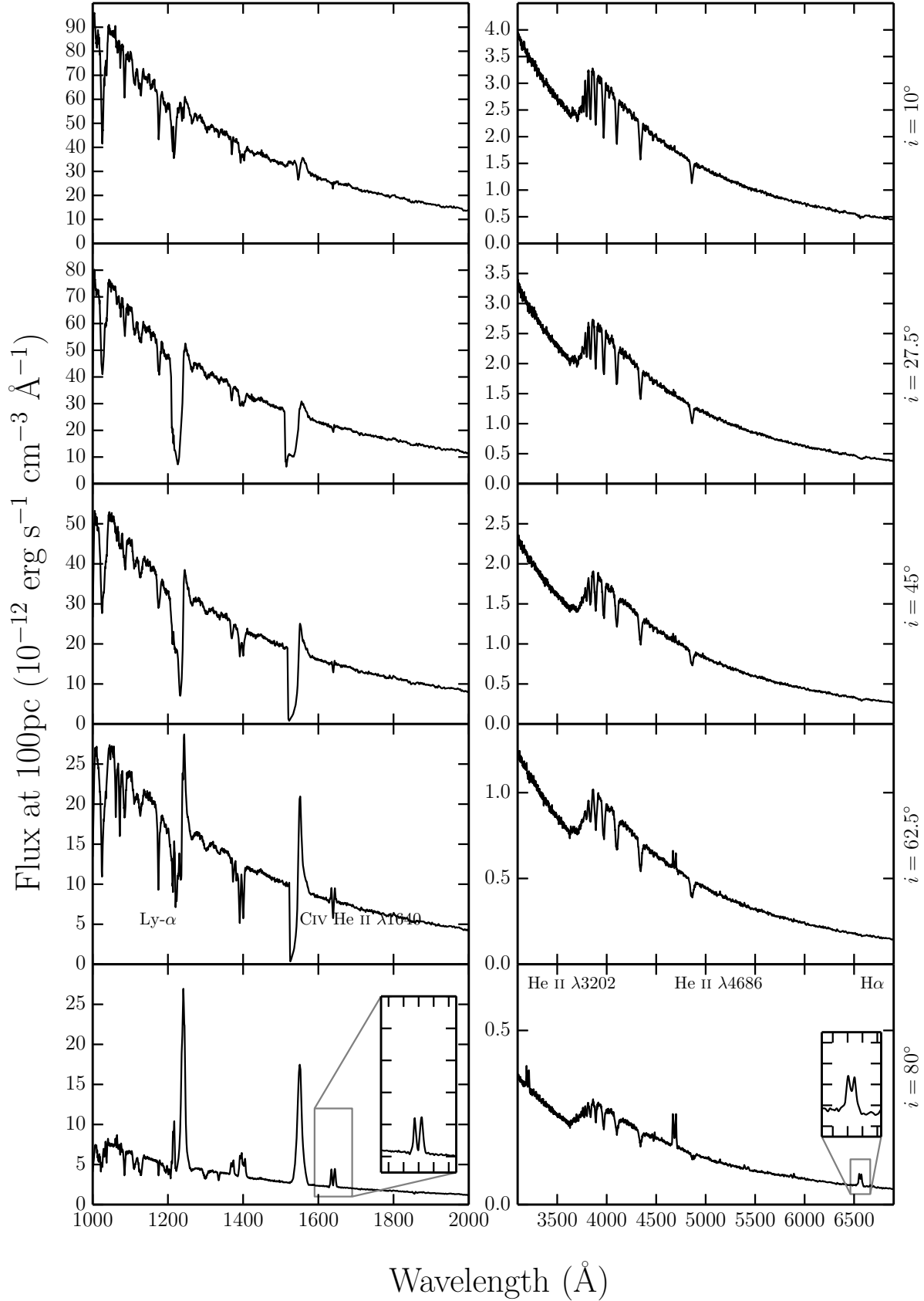


**Figure 8.** Synthetic optical spectra computed for sightlines of 10, 27.5, 45, 62.5 and 80 degrees. In these plots the flux is divided by a polynomial fit to the underlying continuum redward of the Balmer edge, so that line-to-continuum ratios and the true depth of the Balmer jump can be shown.

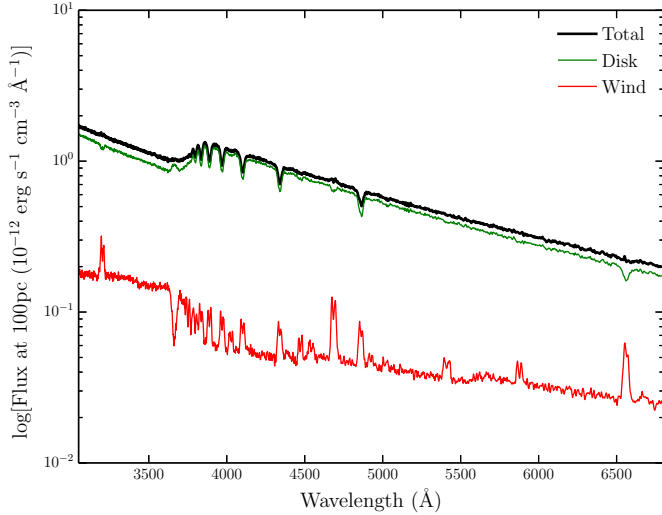
our models is not high enough for this radiative transfer effect to dominate the line shapes.

The Balmer jump is in absorption at all inclinations for our benchmark model. This is due to the stellar atmospheres we have used to model the disk spectrum; it is not a result of photoabsorption in the wind. In fact, the wind spectrum exhibits the Balmer jump in *emission*, but this is not strong enough to overcome the intrinsic absorption edge in the disk spectrum. This is illustrated in Figure 9, which shows the angle-integrated spectrum of the system, i.e. the spectrum formed by all escaping photons, separated into the disk and wind contributions. Even though the wind-formed Balmer recombination continuum does not completely fill in the Balmer absorption edge in this model, it does already contribute significantly to the total spectrum. This suggests that modest changes to the outflow kinematics might boost the wind continuum and produce emergent spectra with weak or absent Balmer absorption edges.





**Figure 7.** UV (left) and optical (right) synthetic spectra for model A, our benchmark model, computed at sightlines of 10, 27.5, 45, 62.5 and 80 degrees. The inset plots show zoomed-in line profiles for He II  $\lambda 1640$  Å and H $\alpha$ . Double-peaked line emission can be seen in He II  $\lambda 1640$  Å, He II  $\lambda 4686$  Å, H $\alpha$  and some He I lines, but the line emission is not always sufficient to overcome the absorption cores from the stellar atmosphere models. The model also produces a prominent He II  $\lambda 3202$  Å line at high inclinations,



**Figure 9.** Total packet-binned spectra across all viewing angles. The thick black line shows the total integrated escaping spectrum, while green and red lines show the contributions from photons that originate in the disk and wind respectively. Recombination continuum emission blueward of the Balmer edge is already prominent relative to other wind continuum processes, but is not sufficient to fill in the Balmer jump in this specific model

## 5 A REVISED MODEL OPTIMIZED FOR OPTICAL WAVELENGTHS

The benchmark model discussed in section 4 was originally designed to reproduce the wind-formed lines seen in the UV spectra of high-state CVs. As we have seen, this model does produce significant amounts of optical emission. However, with only relatively minor changes to the kinematics of the benchmark model, it is possible to construct a model that more closely matches the observed optical spectra of CVs (imaginatively dubbed ‘Model B’).

Specifically, we aim to assess a whether a revised model can:

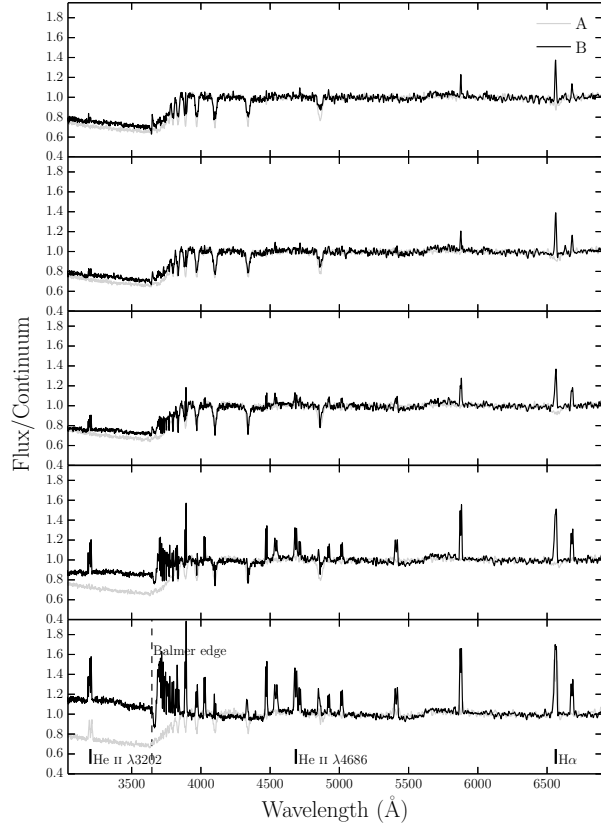
- account for all of the lines we see in optical spectra of CVs while preserving the UV behaviour;
- produce *single-peaked* Balmer emission lines (MC96);
- generate enough of a wind-formed recombination continuum to completely fill in the disk’s Balmer absorption edge for reasonable outflow parameters.

The amount of recombination and collisionally excited emission produced by a given ion,  $i$ , in a plasma of volume  $V$  depends primarily on the emission measure,  $EM$ , defined as

$$EM = \int_0^V n_e n_i dV, \quad (12)$$

where  $n_i$  is the number density of the ion in question. The emission reaching an observer at a given frequency also depends on the Sobolev optical depth it experiences, which depends on both the density and the velocity gradient projected along the line of sight. This can also significantly affect line profile shapes.

The simplest way to simultaneously affect the density in the wind as well as the velocity gradients is by modifying the poloidal velocity law. We have therefore conducted a



**Figure 11.** Synthetic optical spectra computed for sightlines of 10, 27.5, 45, 62.5 and 80 degrees. In these plots the flux is divided by a polynomial fit to the underlying continuum redward of the Balmer edge, so that line-to-continuum ratios and the true depth of the Balmer jump can be shown.

exploratory sensitivity study in which we focus on just two kinematic variables (section 3.1):

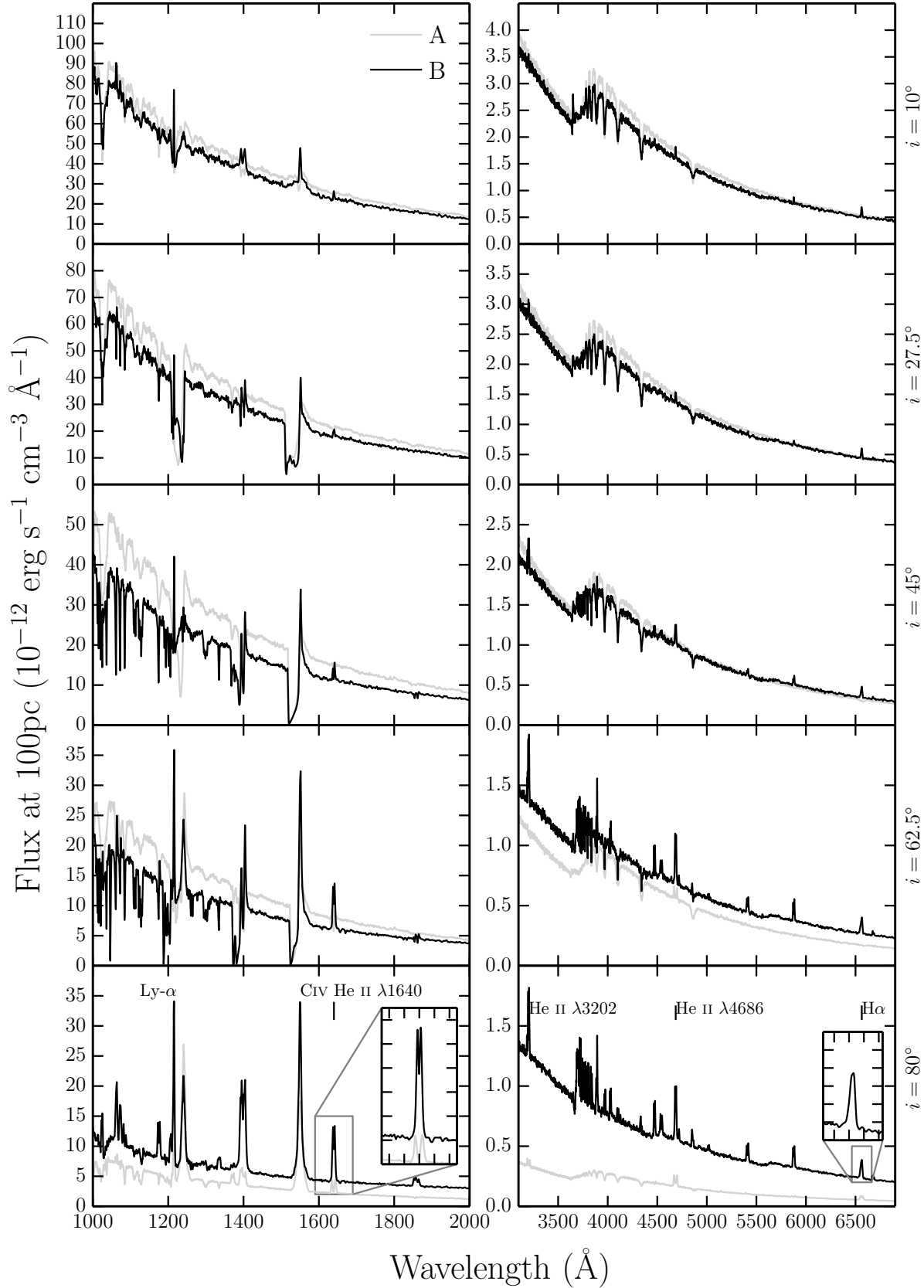
- the acceleration length,  $R_v$ , which controls the distance over which the wind accelerates to  $\frac{1}{2}v_\infty$ ;
- the acceleration exponent,  $\alpha$ , which controls the rate at which the poloidal velocity changes near  $R_v$ .

The general behaviour we might expect is that outflows with denser regions near the wind base - i.e. winds with larger  $R_v$  and/or larger  $\alpha$  - will produce stronger optical emission signatures. However, this behaviour may be moderated by the effect of the increasing optical depth through this region.

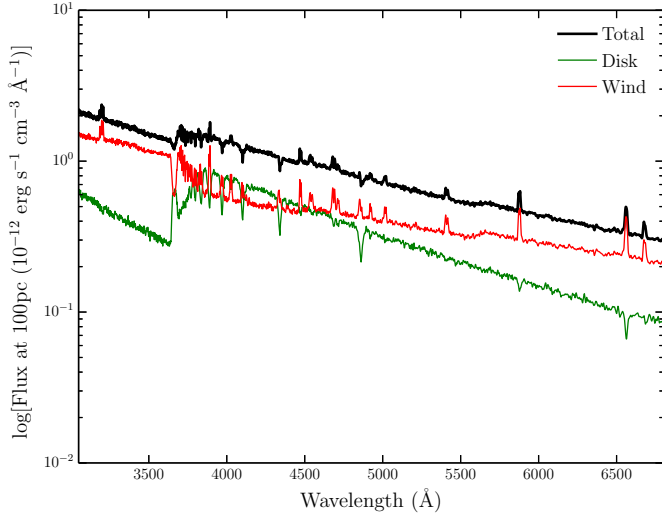
We adopt the parameters listed in table 1 for our model, which best matches the observed optical spectra of high-state CVs. Some of the trends we observe in our parameter search are discussed in section 6. Here, we discuss our revised model.

### 5.1 Synthetic Spectra

Figure 10 shows the predicted UV and optical spectra for this optically optimized model for the full range of inclinations. As expected, the trend from absorption to emission in the optical is again present, but in this improved model we produce emission lines in the entire Balmer series at high in-



**Figure 10.** UV (left) and Optical synthetic spectra for model B computed at sightlines of 10, 27.5, 45, 62.5 and 80 degrees. The inset plots show zoomed-in line profiles for He II  $\lambda 1640$  Å and H $\alpha$ . The H $\alpha$  line is single peaked, but higher order lines in the Balmer series are double-peaked, albeit with narrower profiles. Strong He II  $\lambda 4686$  Å emission can be seen, as well a trend of a deeper Balmer jump with decreasing inclination.



**Figure 12.** Total packet-binned spectra across all viewing angles. The thick black line shows the total integrated escaping spectrum, while green and red lines show the contributions from photons that originate in the disk and wind respectively. Recombination continuum emission blueward of the Balmer edge is already prominent relative to other wind continuum processes, but is not sufficient to fill in the Balmer jump in this specific model

inclinations, as well as the observed lines in He II and He I. This can be seen more clearly in the continuum-divided spectrum in figure 11, with the lines marked (NOTE, mark these).

Two other features are worth noting in the optical spectrum. First, the collisionally excited Ca II H&K emission lines at 3969 Å and (especially) 3934 Å become quite prominent in our densest models. Second, our model predicts a detectable He II recombination line at 3202 Å. This is the Helium equivalent of Paschen  $\beta$  and should be expected in all systems that feature a strong He II 4686 Å line (the Helium equivalent of Paschen  $\alpha$ ). This line is somewhat unfamiliar observationally, because it lies bluewards of the atmospheric cut-off, but also redwards of most ultraviolet spectra.

In the UV, the model still produces all the observed lines, and deep P-Cygni profiles are produced in the normal resonance lines, as discussed in section 4.2. However, the UV spectra also display what is perhaps the biggest problem with this revised model, namely the strength of resonance line emission at low and intermediate inclinations. In order to generate strong optical wind signatures, we have adopted wind parameters that lead to very high densities at the base of the wind ( $n_e \sim 10^{13} - 10^{14}$  in the main emission region). This produces the desired optical recombination emission, but also increases the role of collisional excitation in the formation of the UV resonance lines. This explains the pronounced increase in the emission component of the C IV 1550 Å resonance line, for example, relative to what was seen in the benchmark model (compare figures 7 and 10). The strength of this component in the revised model is probably somewhat too high to be consistent with UV observations of high-state CVs (see e.g. Long et al. 1991, 1994; Noebauer et al. 2010).

## 5.2 Continuum Shape and the Balmer Jump

The wind now also has a clear effect on the continuum shape, as shown by figure 12. The majority of the escaping spectrum has now been reprocessed in some way by the wind, either by electron scattering (the wind is now moderately Thompson-thick), or by bound-free processes. This is demonstrated by the flatter spectral shape and the slight Helium photoabsorption edges present in the optical spectrum. In addition, figures 11 and 12 clearly demonstrate that the wind produces a recombination continuum sufficient to completely fill in the Balmer jump at high inclinations.<sup>2</sup> This offers a potential solution to a long-standing problem in understanding the spectral energy distribution of CVs, as originally suggested by Knigge et al. (1998b).

It has been noted by La Dous (1989b,a) that the Balmer jump in high-state CVs would naturally weaken at high inclinations due to limb darkening effects. Although we include a simple limb darkening law which affects the seen at each inclination, we cannot include it as a *frequency dependent* opacity in our model. As a result, the efficiency of filling in the Balmer jump should really be judged at low and medium inclinations, where, although prominent, the recombination continuum does not overcome the disk atmosphere absorption. In addition, this effect could mean that any model which successfully fills in the jump at low inclinations could lead to a Balmer jump in emission at high inclinations. To understand this phenomenon, a fully self-consistent radiative transfer calculation of both the disk atmosphere and connected wind is required.

## 5.3 Line Profile Shapes: Producing Single-Peaked Emission

Figure 13 shows how the H $\alpha$  line changes with the kinematics of the wind for an inclination of 80°. The line narrows as the models move towards denser, more slowly accelerating winds. This is due the base of the wind becoming optically thick such that line emission cannot escape. This means what is effectively a photosphere for the line moves to a greater vertical distance from the disk plane, where rotational velocities drop off as  $1/r$ . The narrowing of the line is *not* due to the radial velocity shear mechanism proposed by MC96. The main criterion for this mechanism is

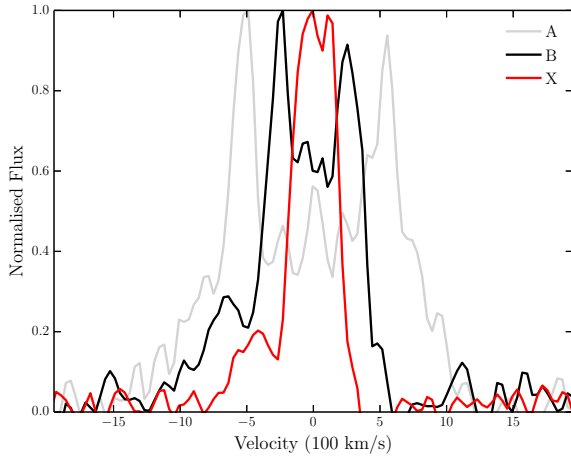
$$\frac{dv_p}{dr} > \frac{dv_\phi}{ds} \quad (13)$$

where  $v_p$  is the poloidal velocity and  $v_\phi$  is the Keplerian velocity. This criterion is simply not satisfied in the emission region of our models.

There are two possible ways in which the model could be altered to better reproduce single-peaked lines as a result of radial velocity shear. First, one could adopt a more equatorial wind geometry. This would increase the velocity gradient in the radial direction. However, CV winds are known to be moderately collimated. This is demonstrated by the fact

<sup>2</sup> Note that the apparent absorption feature just redward of the Balmer jump in these models is artificial. It is caused by residual line blanketing in the stellar atmospheres, which our models cannot fill in since they employ a 20-level Hydrogen atom.





**Figure 13.**  $H\alpha$  line profiles, normalised to 1, plotted in velocity space for three models with varying kinematic properties, computed at an inclination of  $80^\circ$ . The benchmark model and the improved optical model described in section 6 are labeled as A and B respectively, and a third model (X) which has an increased acceleration length of  $R_V = 283.8 R_{WD}$ , and  $\alpha = 4$  is also shown. The  $x$ -axis limits correspond to the Keplerian velocity at  $4R_{WD}$ , the inner edge of the wind. We observe a narrowing of the lines but this is not due to the radial velocity shear mechanism (see section 5.3 for a discussion).

that blue-shifted resonance line absorption is seen in low inclination systems such as RW Sex (Greenstein & Oke 1982), but not in eclipsing systems (Cordova & Mason 1985). Furthermore, the main reason for adopting this type of wind geometry is that it already successfully models the P-Cygni resonance lines seen in the UV (LK02, Noebauer et al. 2010). We note that equatorial winds are expected in AGN and X-ray binaries (Ponti et al. 2012), and that these winds may be more likely to exhibit the required kinematics. Second, one could adopt velocity laws with higher velocity gradients in the line emission region. However, this poses problems for a smooth biconical wind. We have already demonstrated that high densities are required to produce significant emission. In a smooth model, which satisfies the mass continuity condition (see section 3.1), it is therefore difficult to reconcile these high densities with the high velocity gradients required.

We suggest that clumping in CV winds provides a possible solution. Non-homogeneity could allow higher density pockets of emission to exist with radial velocity shear still being high. Line-driven winds are expected to be unstable (REFS), and a number of implementations have been used in the stellar wind community (REFS). We therefore hope that further investigation of clumped winds could help solve the current ‘mystery’ of single-peaked line emission in high inclination systems.

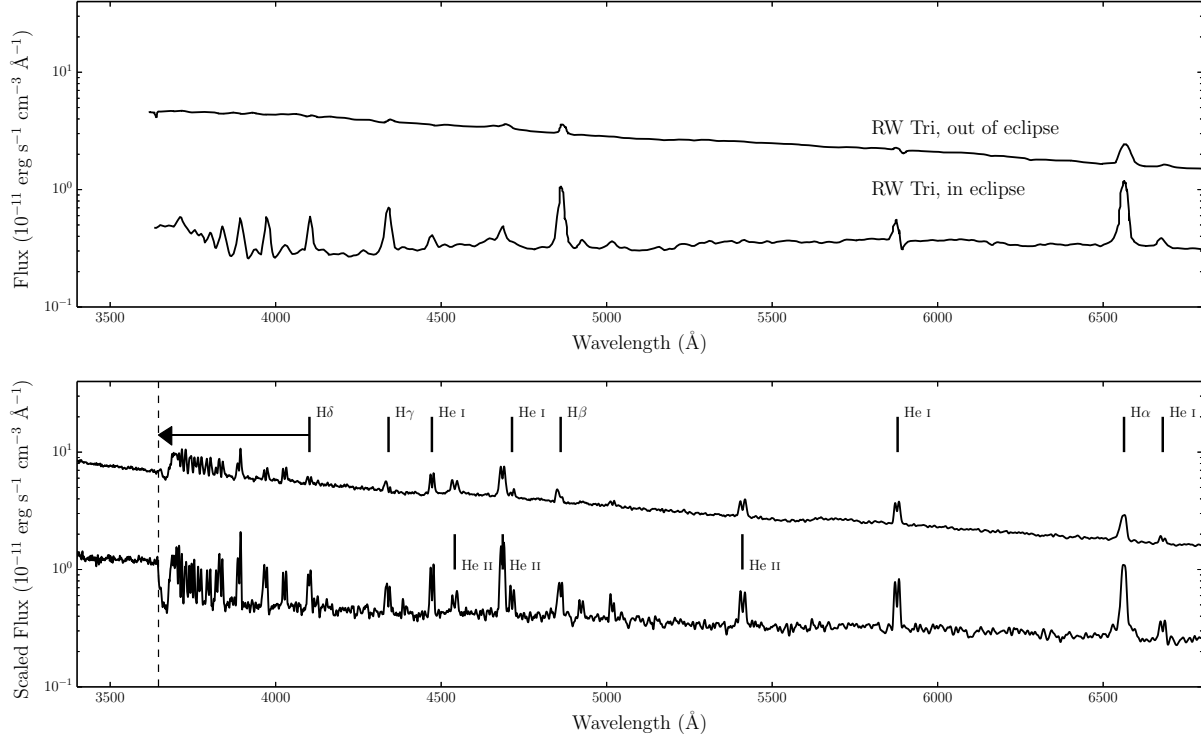
#### 5.4 Comparison to RW Tri

Figure 14 shows a qualitative comparison of the predicted out-of-eclipse and mid-eclipse spectra against observations of the high-inclination nova-like RW Tri. The inclination of RW Tri is somewhat uncertain, with estimates including

$70.5^\circ$  (Smak 1995),  $75^\circ$  (Groot et al. 2004),  $80^\circ$  (Longmore et al. 1981) and  $82^\circ$  (Frank & King 1981). Here, we adopt  $i = 80^\circ$ , but our qualitative conclusions are not particularly sensitive to this particular value. We emphasize that this model is in no sense a fit to this - or any other - data set.

The similarity between the synthetic and observed spectra is striking. In particular, the revised model produces strong emission in all the Balmer lines, with line-to-continuum ratios comparable to those seen in RW Tri. Moreover, the line-to-continuum contrast increases during eclipse, as expected for emission produced in a disk wind. This trend is in line with the observations of RW Tri, and it has also been seen in other NLs, including members of the SW Sex class (Neustroev et al. 2011).

There are also interesting differences between the revised model and the RW Tri data set. For example, the model exhibits considerably stronger He II features than the observations, which suggests that the overall ionization state of the model is somewhat too high. As discussed in section 5.3, the optical lines are narrow. This is in contrast to what is generally seen in observations of NLs, although the relatively low resolution of the RW Tri spectra makes a specific comparison difficult.



**Figure 14.** *Top Panel:* In and out of eclipse spectra of the high inclination NL RW Tri. *Bottom Panel:* In and out of eclipse synthetic spectra from Model B. The artificial ‘absorption’ feature just redward of the Balmer jump is caused for the reasons described in section 5.2.

## 6 CONCLUSIONS

We have investigated whether a disk wind model designed to reproduce the UV spectra of high-state CVs would also have a significant effect on the optical spectra of these systems. We find that this is indeed the case. In particular, the model wind produces H and He recombination lines, as well as a recombination continuum blueward of the Balmer edge. We do not produce P-Cygni profiles in the optical H and He lines, which are seen in a small fraction of CV optical spectra. Possible reasons for this are briefly discussed in section 4.2.

We have also constructed a revised benchmark model which is designed to more closely match the optical spectra of high-state CVs. This optically optimized model produces all the prominent optical lines in and out of eclipse, and achieves reasonable verisimilitude with the observed optical spectra of RW Tri. However, this model also has significant shortcomings. In particular, it predicts stronger-than-observed He II lines in the optical region and too much of a collisionally excited contribution to the UV resonance lines that

Based on this, we show that recombination emission from outflows with sufficiently high densities and/or optical depths might

- fill in the Balmer absorption edge in the spectrum of the accretion, thus accounting for the absence of a strong edge in observed CV spectra (Knigge et al. 1998b);
- produce the optical lines observed in CVs.

Although we note that limb darkening effects on disk

atmosphere opacities would affect the Balmer jump (see section 5.2). In section 5.3, we demonstrate that although single-peaked emission can be formed in our densest models, it is not due to the radial velocity shear mechanism proposed by MC96. The emission in our models which produce single-peaked lines comes from approximately  $10^{11}$  cm above the disk plane, which is not supported by eclipse-mapping observations (REFS). We suggest that non-homogenous, ‘clumpy’ line-driven winds may offer a potential solution.

We can also make a number of basic, testable predictions for wind-formed signatures:

- we expect the depth of the Balmer absorption edge in CV spectra to decrease with increasing inclination (with the Balmer jump possibly switching to emission at the highest inclinations, if the flux in the wind-formed recombination continuum exceeds that of the foreshortened disk spectrum);
- we predict that a strong He II  $\lambda 3202\text{\AA}$  line should be prominent in all cases where He II  $\lambda 1640\text{\AA}$  & He II  $\lambda 4686\text{\AA}$  are seen (but note that this prediction is not restricted to lines formed in disk winds);
- we predict that the Balmer jump should be in emission during mid-eclipse in high inclination systems.

Overall, our work demonstrates that *disk winds matter*. They are not just responsible for creating the blue-shifted absorption and P-Cygni profiles seen in the UV resonance lines of high-state CVs, but can also have a strong effect on the optical appearance of these systems. In fact, most of the optical features characteristic of CVs are likely to be affected – and possibly even dominated – by their disk winds. Given that optical spectroscopy plays the central role in observa-

tional studies of CVs, it is clearly critical to know where and how these spectra are actually formed. We believe it is high time for a renewed effort to understand the formation of spectra in accretion disks and associated outflows.

## Acknowledgements

The work of JHM and CK is supported by the Science and Technology Facilities Council (STFC), via studentships and a consolidated grant, respectively. We would like to thank Juan Hernandez Santisteban, Samuel Mangham and Ivan Hubeny for useful discussions. We acknowledge the use of the IRIDIS High Performance Computing Facility, and associated support services at the University of Southampton, in the completion of this work.

## REFERENCES

- Badnell N. R., Bautista M. A., Butler K., Delahaye F., Mendoza C., Palmeri P., Zeppen C. J., Seaton M. J., 2005, *MNRAS* 360, 458
- Beuermann K., Stasiewski U., Schwobe A. D., 1992, *A&A* 256, 433
- Beuermann K., Thomas H.-C., 1990, *A&A* 230, 326
- Cordova F. A., Mason K. O., 1982, *ApJ* 260, 716
- Cordova F. A., Mason K. O., 1985, *ApJ* 290, 671
- Dhillon V. S., 1996, in A. Evans, J. H. Wood (eds.), *IAU Colloq. 158: Cataclysmic Variables and Related Objects*, Vol. 208 of *Astrophysics and Space Science Library*, 3
- Dhillon V. S., Rutten R. G. M., 1995, *MNRAS* 277, 777
- Drew J., Verbunt F., 1985, *MNRAS* 213, 191
- Ferland G. J., Porter R. L., van Hoof P. A. M., Williams R. J. R., Abel N. P., Lykins M. L., Shaw G., Henney W. J., Stancil P. C., 2013, *RMxAA* 49, 137
- Frank J., King A. R., 1981, *MNRAS* 195, 227
- Greenstein J. L., Oke J. B., 1982, *ApJ* 258, 209
- Groot P. J., Rutten R. G. M., van Paradijs J., 2004, *A&A* 417, 283
- Haug K., 1987, *AP&SS* 130, 91
- Heap S. R., Boggess A., Holm A., Klingle-Smith D. A., Sparks W., West D., Wu C. C., Boksenberg A., Willis A., Wilson R., Macchetto F., Selvelli P. O., Stickland D., Greenstein J. L., Hutchings J. B., Underhill A. B., Viotti R., Whelan J. A. J., 1978, *Nature* 275, 385
- Higginbottom N., Knigge C., Long K. S., Sim S. A., Matthews J. H., 2013, *MNRAS* 436, 1390
- Honeycutt R. K., Schlegel E. M., Kaitchuck R. H., 1986, *ApJ* 302, 388
- Horne K., Marsh T. R., 1986, *MNRAS* 218, 761
- Hubeny I., Lanz T., 1995, *ApJ* 439, 875
- Idan I., Lasota J.-P., Hameury J.-M., Shaviv G., 2010, *A&A* 519, A117
- Kafka S., Honeycutt R. K., 2004, *AJ* 128, 2420
- Kasen D., Thomas R. C., Nugent P., 2006, *ApJ* 651, 366
- Kerzendorf W. E., Sim S. A., 2014, *MNRAS* 440, 387
- Knigge C., Drew J. E., 1997, *ApJ* 486, 445
- Knigge C., Long K. S., Blair W. P., Wade R. A., 1997, *ApJ* 476, 291
- Knigge C., Long K. S., Wade R. A., Baptista R., Horne K., Hubeny I., Rutten R. G. M., 1998a, *ApJ* 499, 414
- Knigge C., Long K. S., Wade R. A., Baptista R., Horne K., Hubeny I., Rutten R. G. M., 1998b, *ApJ* 499, 414
- Knigge C., Woods J. A., Drew J. E., 1995, *MNRAS* 273, 225
- Kurucz R. L., 1991, in L. Crivellari, I. Hubeny, D. G. Hummer (eds.), *NATO ASIC Proc. 341: Stellar Atmospheres - Beyond Classical Models*, 441
- La Dous C., 1989a, *MNRAS* 238, 935
- La Dous C., 1989b, *A&A* 211, 131
- Long K. S., Blair W. P., Davidsen A. F., Bowers C. W., Dixon W. V. D., Durrance S. T., Feldman P. D., Henry R. C., Kriss G. A., Kruk J. W., Moos H. W., Vancura O., Ferguson H. C., Kimble R. A., 1991, *ApJ Letters* 381, L25
- Long K. S., Knigge C., 2002, *ApJ* 579, 725
- Long K. S., Wade R. A., Blair W. P., Davidsen A. F., Hubeny I., 1994, *ApJ* 426, 704
- Longmore A. J., Lee T. J., Allen D. A., Adams D. J., 1981, *MNRAS* 195, 825
- Lucy L. B., 2002, *A&A* 384, 725
- Lucy L. B., 2003, *A&A* 403, 261
- Marsh T. R., Horne K., 1990, *ApJ* 349, 593
- Mauche C. W., Raymond J. C., 1987, *ApJ* 323, 690
- Mihalas D. M., 1982, *Stellar atmospheres*.
- Murray N., Chiang J., 1996, *Nature* 382, 789
- Murray N., Chiang J., 1997, *ApJ* 474, 91
- Neustroev V. V., Suleimanov V. F., Borisov N. V., Belyakov K. V., Shearer A., 2011, *MNRAS* 410, 963
- Noebauer U. M., Long K. S., Sim S. A., Knigge C., 2010, *ApJ* 719, 1932
- Osterbrock D. E., 1989, *Astrophysics of gaseous nebulae and active galactic nuclei*
- Patterson J., Patino R., Thorstensen J. R., Harvey D., Skillman D. R., Ringwald F. A., 1996, *AJ* 111, 2422
- Ponti G., Fender R. P., Begelman M. C., Dunn R. J. H., Neilsen J., Coriat M., 2012, *MNRAS* 422, L11
- Puebla R. E., Diaz M. P., Hillier D. J., Hubeny I., 2011, *ApJ* 736, 17
- Ringwald F. A., Naylor T., 1998, *AJ* 115, 286
- Seaton M. J., 1959, *MNRAS* 119, 90
- Shakura N. I., Sunyaev R. A., 1973, *A&A* 24, 337
- Shaviv G., Wehrse R., 1991, *A&A* 251, 117
- Shlosman I., Vitello P., 1993, *ApJ* 409, 372
- Sim S. A., 2004, *MNRAS* 349, 899
- Sim S. A., Drew J. E., Long K. S., 2005, *MNRAS* 363, 615
- Sim S. A., Long K. S., Miller L., Turner T. J., 2008, *MNRAS* 388, 611
- Smak J., 1981, *ACTAA* 31, 395
- Smak J., 1995, *ACTAA* 45, 259
- Tatum M. M., Turner T. J., Sim S. A., Miller L., Reeves J. N., Patrick A. R., Long K. S., 2012, *ApJ* 752, 94
- van Regemorter H., 1962, *ApJ* 136, 906
- Wade R. A., 1984, *MNRAS* 208, 381
- Wade R. A., 1988, *ApJ* 335, 394