

# Testing Quasar Unification with Clumpy Wind Models

James H. Matthews<sup>1\*</sup>, Christian Knigge<sup>1</sup>, Nick Higginbottom<sup>1</sup>, Knox S. Long<sup>2</sup>,  
Stuart A. Sim<sup>3</sup> and Sam W. Mangham<sup>1</sup>

<sup>1</sup>*School of Physics and Astronomy, University of Southampton, Highfield, Southampton, SO17 1BJ, United Kingdom*

<sup>2</sup>*Space Telescope Science Institute, 3700 San Martin Drive, Baltimore, MD, 21218*

<sup>3</sup>*School of Mathematics and Physics, Queens University Belfast, University Road, Belfast, BT7 1NN, Northern Ireland, UK*

19 September 2015

## ABSTRACT

Various unification schemes have been proposed to interpret the complex phenomenology of quasars and luminous active galactic nuclei (AGN) in terms of a simple axisymmetric picture involving a central black hole, an accretion disc and an associated outflow. Here, we continue our tests of this paradigm by comparing “average” AGN spectra to synthetic spectra of simple biconical disc wind models, using our state-of-the-art Monte Carlo radiative transfer and photoionization code. In particular, we investigate whether clumping of the outflow allows us to produce synthetic spectra in the rest-frame UV that have the characteristics of AGN, assuming total luminosities as high as  $10^{45}$  ergs s<sup>-1</sup>. We find that a simple treatment of clumping (‘microclumping’) allows for a more realistic AGN luminosities, while maintaining the ionization state necessary for strong BAL features in the rest-frame UV. We examine the X-ray properties of these simple clumped models and find good agreement with existing X-ray samples of AGN and quasars. The dense, X-ray heated wind produces strong recombination and collisionally excited line emission in, e.g., C IV and Ly  $\alpha$ , to emerge at the low inclination, ‘Type 1 quasar-like’ angles. At the highest inclinations, the synthetic spectra possess prominent Mg II and Al III BALs, the absorption features seen in LoBAL quasars. Despite these successes, we are unable to reproduce the remarkably uniform emission line properties seen in BAL and non-BAL quasar composites. **This sentence needs to be more precise, especially in light of the statement earlier in the abstract about “good agreement”** This is due to a fundamental constraint arising from the anisotropy of emission from a classical thin disc. Overall, our work suggests that geometric unification involving an accretion disc wind is a promising scenario, but our results pose a number of difficult challenges to such a model. **There is confusion in the abstract, at least, in what is the model and what is the synthetic spectra, as well as whether there is one model or many models. We need to try to be clear. The model should be a clumped biconical outflow. There are then particular instantiations of the model. And then there are synthetic spectra**

## 1 INTRODUCTION

The spectra of quasars and luminous active galactic nuclei (AGN) typically exhibit a series of strong emission lines with an underlying blue continuum - the so-called ‘*big blue bump*’ (BBB). The BBB is often attributed to emission from a geometrically thin, optically thick accretion disc surrounding the central black hole, similar to that described by Shakura & Sunyaev (1973). The lines are usually attributed to scattering and absorption of material above the disk. **Note: the first half of this paragraph and the second do not follow from one another. The first half is about spectra the second about outflows and the connection to spectra is not made until the next paragraph. This needs fixing. JM: Knox, I like this structure...I start with inflow, then move to outflow.** In addition to the in-

flowing accreting material, outflows are ubiquitous in AGN and quasars (Kellermann et al. 1989; Ganguly & Brotherton 2008). These outflows can take the form of highly collimated radio jets (e.g. Hazard et al. 1963; Potash & Wardle 1980; Perley et al. 1984; Marscher 2006), or mass-loaded ‘winds’ emanating from the accretion disc (Weymann et al. 1991; Turner & Miller 2009). Outflows in AGN offer a potential feedback mechanism through which the central source can affect its environment (King 2003, 2005; Fabian 2012) – feedback that is required in models of galaxy evolution (Springel et al. 2005) and may explain the ‘ $M - \sigma$ ’ relation (Silk & Rees 1998; Häring & Rix 2004).

Approximately 20% of quasars exhibit blueshifted ( $\sim 0.1c$ ) broad absorption lines (BALs) in the ultraviolet, providing clear evidence for outflowing absorbing material

(Weymann et al. 1991; Reichard et al. 2003; Knigge et al. 2008; Turner & Miller 2009; Allen et al. 2011). The simplest explanation for the incidence of BAL quasars (BALQSOs) is in terms of an accretion disc wind (ADW). Within the ADW paradigm, a biconical wind rises from the accretion disc and the BALQSO fraction is associated with the covering factor of the outflow. Polarisation studies expect the wind to be roughly equatorial (Goodrich & Miller 1995; Cohen et al. 1995), although this geometry is contradicted by more recent studies of radio spectral indices (REFS DIPOMPEO; see section ?). ADWs offer a natural explanation for the diverse phenomenology of luminous AGN and QSOs (e.g. Murray et al. 1995; Elvis 2000). Depending on viewing angle, an observer may see a BALQSO or normal ‘Type 1’ quasar. Within this framework, the broad-line region (BLR) can correspond either to the dense wind base or dense clumps embedded in the outflows. Indeed, Elitzur et al. (2014) show that a disc-wind BLR scenario naturally explains the emission line evolution of AGN. A biconical wind model can also readily explain the various sub-classifications of BALQSOs: HiBALQSOs, which only exhibit high ionization line absorption; LoBALQSOs, which also show absorption in lower ionization state species such as Mg II and Al III; and FeLoBALQSOs which show further absorption in Fe II and III. In unified geometric models, this is generally attributed to ionization stratification of the outflow (e.g. Elvis 2000).

As well as imprinting clear line absorption and emission features, disc winds may also have a profound effect on the structure and emergent *continuum* of the accretion disc itself. Mass-loss will alter the accretion rate and resultant temperature of the accretion disc, possibly explaining some of the features typically seen in luminous AGN (Knigge 1999; Laor & Davis 2014). There have been numerous difficulties when confronting theoretical accretion disc models with observations (see e.g. Blaes 1998). However, AGN spectral energy distributions (SEDs) can now, in general, be fitted well with accretion disc models when the effects of general relativity (GR), Comptonisation and mass-loss are included (Capellupo et al. 2015). Mass-loss therefore appears to be critical if an accretion disc model is to successfully fit AGN SEDs, particular in the UV region of the spectrum.

**ksl: I don’t see what this paragraph has to do with our problem, which we need to clearly explain**  
Despite the clear importance of ADWs in understanding AGN SEDs and accretion physics, much of the underlying outflow physics remains highly uncertain. Several possible driving mechanisms for ADWs have been proposed, including thermal pressure (Weymann et al. 1982; Begelman et al. 1991), magnetocentrifugal forces (Blandford & Payne 1982; Pelletier & Pudritz 1992) and radiation pressure on spectral lines (‘line-driving’; Lucy & Solomon 1970; Shlosman et al. 1985; Murray et al. 1995). Of these, line-driving is possibly the most attractive, as strong absorption lines are already seen in BALQSOs and the X-ray spectra of AGN (Reeves et al. 2003; Pounds & Reeves 2009; Tombesi et al. 2010). The presence of line-locked features (Bowler et al. 2014) and the ‘ghost of Ly  $\alpha$ ’ (Arav et al. 1996; Arav 1996; North 2006; but see also Cottis et al. 2010) in the spectra of some BALQSOs also gives clearer evidence that line-driving is at least partially contributing to the acceleration of the wind.

The efficiency of line-driving is crucially dependent on the ionization state of the outflowing plasma, meaning that

it is difficult to prevent the wind becoming over-ionized and ‘failing’ in the presence of strong X-rays. Murray et al. (1995) proposed a potential solution: a region of ‘hitchhiking gas’ that could shield the wind from the central X-ray source. Hydrodynamic simulations of line-driven disc winds also found a shielding region was required to maintain the correct ionization state (Proga et al. 2000; Proga & Kallman 2004). However, Higginbottom et al. (2014) showed that including multiple scattering means the ionizing radiation field could still reach the previously shielded regions in those particular models. An additional or alternative solution is that the wind is clumped (e.g. Hamann et al. 2013) possibly on multiple scale lengths. Local density enhancements could lower the ionization parameter of the plasma while still maintaining the same mass-loss rate and column density.

Evidence for dense substructures in AGN winds is widespread. BALQSOs show complex absorption line profiles (Ganguly et al. 2006; Simon & Hamann 2010) and exhibit variability in these profile shapes (Capellupo et al. 2011, 2012, 2014). AGN generally show variability in X-ray absorption components (e.g. Risaliti et al. 2002) and many models for the BLR consist of clumps embedded in an outflow (Krolik et al. 1981; Emmering et al. 1992; de Kool & Begelman 1995; Cassidy & Raine 1996). Clumping can be caused by magnetic confinement de Kool & Begelman (1995), or the instabilities inherent to line-driven winds (Lucy & Solomon 1970; MacGregor et al. 1979; Carlberg 1980; Owocki & Rybicki 1984, 1985). Additionally, clumping is required to explain the electron scattering wings of emission lines formed in line-driven hot star winds (Hillier 1991). Complex substructures on a variety of scales are also produced in simulations of line-driven outflows in AGN (Proga et al. 2000; Proga & Kallman 2004; Proga & Kurosawa 2010; Proga et al. 2014). Clumpy winds therefore offer an observationally motivated and theoretically predicted way to lower the ionization state of a plasma, possibly in tandem with a shielding scenario.

We have been engaged in a project to determine whether it is possible to simulate the properties of the spectra of AGN, including BAL QSOs, using simple kinematic prescriptions for biconical disc winds using a Monte Carlo radiative transfer code that calculates the ionization structure of the wind and simulates the spectra from such a system (Sim et al. 2008, 2010; Higginbottom et al. 2013, hereafter H13). The results have encouraging in the sense that in H13, we showed we could produce simulated spectra that resembled that of BAL QSOs, as long as the luminosity of the central engine was relatively low, of order  $10^{43}$  ergs s<sup>-1</sup> and the mass loss rate was relatively high, of order the mass accretion rate. However, at higher luminosities, the wind was so ionized that UV absorption lines were not produced. In addition, and in part due to limitations in our radiative transfer code, the model failed to produce spectra with strong emission lines at any inclination angle.

Here we attempt to address both of these issues, by introducing clumping into our model and a more complete treatment of H and He into our radiative transfer calculations. The remainder of this paper is organized as follows: In section 2, we describe some of the important photoionization and MCRT aspects of the code. In section 3, we outline the model, including a description of our clumping implemen-

tation. In section 4, we present the results from a clumped model. In section 5 we discuss our results, focussing particularly on the anisotropy of disc emission and GR effects, and finally, in section 6, we summarise our findings.

## 2 IONIZATION AND RADIATIVE TRANSFER

For this study, we use the MCRT code PYTHON we have developed to carry out our radiative transfer and photoionization calculations in non-local thermodynamic equilibrium (non-LTE) in a variety of disk-wind systems. We have used the code to model accreting white dwarfs (Long & Knigge 2002, hereafter LK02; Noebauer et al. 2010; Matthews et al. 2015, hereafter M15), young-stellar objects (Sim et al. 2005) and quasars (H13). LK02 provides the basic description of the original code; improvements, see especially H14 and M15, have been described in subsequent reports. A full description of the code will be published elsewhere; here we focus on the changes made to improve the ionization calculation and to allow for clumping in the wind.

**ksl - I a a bit torn over whether we add a few 3 sentences to describe the basic idea. Some will be reading this for the first time. Items that one would include are that we calculate the radiative transfer in a cylindrical coordinate system of dimension ?? x ?? . We have ionization and spectral cycles. In each cycle, ?? photon bundles are generated. We stop the ionization cycles after they converge, without any details. I think this would be useful because we introduce terms like “cell” later on.’, and without some kind of description they have no context**

### 2.1 Line transfer

Our approach to line transfer is based upon the macro-atom implementation developed by Lucy (2002, 2003). In our case, for reasons of computational efficiency, we adopt the hybrid macro-atom scheme described by M15 (an improvement over that used by H14 for our earlier quasar study), in which the energy flows through the system are described in terms of indivisible energy quanta of radiant or kinetic energy (‘*r*–packets’ and ‘*k*–packets’ respectively; see also section 2.3). These energy packets interact with either two-level ‘simple ions’ or full ‘macro-atoms’. Our scheme allows one to treat non-LTE line transfer in radiative equilibrium without approximation for elements which are identified as full macro-atoms, while maintaining the fast ‘two-level’ treatment of resonance lines when elements are identified as simple-ions (see M15). In this study, only H is treated as a macro-atom, because we expect recombination to be important in determining their level populations and resultant line emission, and because we were especially interested in the contribution to AGN spectra of Lyman  $\alpha$ . In H14, all atoms were treated in a two-level approximation.

### 2.2 Ionization Scheme

Macro-atoms have their ion and level populations derived from MC rate estimators as described by M15. Previously (LK02, H13, M15), we used a modified Saha approach to calculate the ionization fractions of simple-ions. As part of this effort, we have now improved PYTHON to explicitly solve

the rate equations between ions in non-LTE. This dispenses with a number of small assumptions made in the modified Saha approach, is more numerically stable, and, in principle, allows the direct addition of extra physical processes such as Auger ionization.

**ksl - Do we need the rest of this, or could it be postponed until a detailed paper on the program**

We model the SED in a grid cell using the technique described by H13. In this scheme, the mean intensity in a series of  $n$  bands is modeled as a normalised power law in frequency  $\nu$

$$J_{\nu,i} = K_{pl} \nu^{\alpha_{pl}}, \quad (1)$$

for a band  $i$ , or an exponential

$$J_{\nu,i} = K_{exp} e^{(-h\nu/kT_{exp})}. \quad (2)$$

Here,  $K_{exp}$ ,  $K_{pl}$ ,  $T_{exp}$  and  $\alpha_{pl}$  are spectral fit parameters deduced from the band-limited radiation field estimators. The ionization rate out of ion  $j$  can then be written as

$$R_{j,j+1}(J) = n_i \left( C_j n_e + \sum_{band} \sum_{i=0}^n \int_{\nu_i}^{\nu_{i+1}} \frac{4\pi J_{\nu,i} \sigma_j(\nu)}{h\nu} d\nu \right), \quad (3)$$

where  $\sigma_j$  is the photoionization cross-section and  $C_j$  represents the collisional ionization coefficient. The recombination rate into ion  $j$  is given by

$$R_{j+1,j}(T_e) = (\alpha_{RR}^j + \alpha_{DR}^j + \alpha_{CR}^j) n_{j+1} n_e, \quad (4)$$

where each  $\alpha^j$  here is the recombination rate coefficient into the ground state of ion  $j$ . The subscripts denote radiative, dielectronic and collisional (three-body) recombination, respectively. For simple-ions, we use a dilute Boltzmann equation to calculate the population of level  $k$  in ionic stage  $j$ ,

$$\frac{n_{jk}}{n_j} = \frac{W g_k}{z_j(T_R)} \exp(-E_k/kT_R). \quad (5)$$

Here  $z_j$  is the partition function of ionic stage  $j$ ,  $T_R$  is the effective radiation temperature,  $E_k$  is the energy difference between level  $k$  and the ground state, and  $g_k$  is the statistical weight of level  $k$ .

### 2.3 Physical Processes

We include all free-free, bound-free and bound-bound heating and cooling processes in the model. For radiative transfer purposes we treat electron scattering in the Thomson limit, but take full account of Compton heating and cooling when calculating the thermal balance of the plasma (see H13). Adiabatic cooling is included and represents the only departure from strict radiative equilibrium, but is insignificant in most of the outflow.

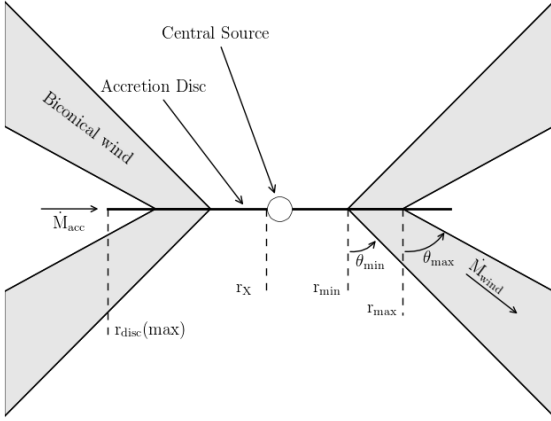
### 2.4 Atomic Data

We use the same atomic data as described by LK02 and since updated by H13 and M15, with the addition of direct ionization data from Dere (2007). Photoionization cross-sections are from TOPBASE (Cunto et al. 1993) and Verner et al. (1996). Dielectronic and radiative recombination rate coefficients are taken from the CHIANTI database version 7.0 (Dere et al. 1997; Landi et al. 2012). We use ground state recombination rates from Badnell (2006) where available, and otherwise default to calculating recombination rates from

the Milne relation. Free-free Gaunt factors are from Sutherland (1998).

### 3 A CLUMPY BICONICAL DISK WIND MODEL FOR QUASARS

Our kinematic prescription for a biconical disc wind model follows Shlosman & Vitello (1993), and is described further by LK02, H13 and M15. A schematic is shown in figure 1, with key aspects marked. The general biconical geometry is similar to that invoked by Murray et al. (1995) and Elvis (2000) in order to explain the phenomenology of quasars and BALQSOs.



**Figure 1.** A cartoon showing the geometry and some key parameters of our biconical wind model.

#### 3.1 Photon Sources

The accretion disc in our model is geometrically thin, but optically thick. We thus treat the disc as an ensemble of blackbodies with a Shakura & Sunyaev (1973) effective temperature profile. The emergent SED is then determined by the specified accretion rate ( $\dot{m}$ ) and central BH mass ( $M_{BH}$ ). All photon sources in our model are assumed to be opaque, meaning that photons which strike them are destroyed. The inner radius of the disc extends to the innermost stable circular orbit (ISCO) of the BH. We assume a Schwarzschild BH with an ISCO at  $6 r_G$ , where  $r_G = GM_{BH}/c^2$  is the gravitational radius. For a  $10^9 M_\odot$  black hole, this is equal to  $8.8 \times 10^{14}$  cm or  $\sim 10^{-4}$  pc.

The X-ray source in our model is treated as an isotropic sphere at the ISCO, which emits r-packets according to a power law in flux with index  $\alpha_X$ , of the form

$$F_X(\nu) = K\nu^{\alpha_X}. \quad (6)$$

The normalisation,  $K_X$  of this power law is such that it produces the specified 2-10 keV luminosity,  $L_X$ . In addition to the disc and X-ray source, the wind is able to reprocess radiation. However, new photon packets are not produced in the wind (as in LK02). Instead, this reprocessing is dealt with by enforcing strict radiative equilibrium (*modulo* adiabatic cooling; see section 2.3) via an indivisible energy packet constraint (see Lucy 2002, M15).

#### 3.2 Kinematics and Geometry

In our model, a biconical disc wind rises from the accretion disc between launch radii  $r_{min}$  and  $r_{max}$ . The opening angles of the wind are set to  $\theta_{min}$  and  $\theta_{max}$ . The poloidal velocity along each individual streamline at a poloidal distance  $l$  is then given by

$$v_l = v_0 + [v_\infty(r_0) - v_0] \frac{(l/R_v)^\alpha}{(l/R_v)^\alpha + 1}, \quad (7)$$

where  $v_0$  is the velocity at the base of the streamline,  $\alpha$  is an exponent governing how quickly the wind accelerations and  $R_v$  is the ‘acceleration length’, defined as the distance at which the outflow reaches half of its terminal velocity,  $v_\infty$ . The terminal velocity is set to a fixed multiple of the escape velocity,  $v_{esc}$ , at the base of the streamline (radius  $r_0$ ). The rotational velocity,  $v_\phi$ , is initially Keplerian ( $v_k = [GM/r_0]^{1/2}$ ), and the wind conserves specific angular momentum, such that

$$v_\phi r = v_k r_0. \quad (8)$$

The velocity law is crucial in determining the output spectra, as it affects not only the projected velocities along the line of sight, but also the density and ionization state of the outflow. A wind which accelerates more slowly will have a denser wind base with correspondingly different ionization and emission characteristics.

#### 3.3 Clumping

Our previous modelling efforts assumed a smooth outflow, in which the density at a given point was determined only by the kinematic parameters and mass loss rate. However, as already discussed, AGN winds exhibit significant substructure – the outflow is expected to be *clumpy*, rather than smooth, and probably on a variety of scales. Implementing a treatment of clumping is challenging, for two main reasons. First, there are significant computational difficulties associated with adequately resolving and realistically modelling a series of small scale, high density regions with a MCRT code. Second, the addition of multiple additional degrees of freedom in the model results in significantly wider parameter space. Unfortunately, the physical scale lengths and density contrasts associated with these parameters cannot be well-constrained from observations.

To allow for clumping in our outflow we adopt a simple approximation used extensively in stellar wind modelling, known as *microclumping* (Hamann & Koesterke 1998)(MORE REFS). The key assumption here is that typical clump sizes are much smaller than the typical photon mean free path, and thus the clumps are both geometrically and optically thin. This approach allows one to introduce a ‘filling factor’,  $f$ , which is the fraction of the volume of the plasma filled by clumps. The intra-clump medium is assumed to be a vacuum. We can then introduce the density enhancement,  $D$ , which is defined as

$$D = \frac{1}{f}. \quad (9)$$

We then multiply all densities in the model by  $D$ , and all emitting volumes by  $f$ . This has the effect of enhancing all emissivities and opacities that scale with the square of density (such as collisional excitation and recombination) by a

factor  $D$ . All processes that scale linearly with density (such as electron scattering and bound-free opacity) will remain unchanged for a given ionization state.

Clumping the wind has an important effect on the ionization state and has been proposed as a solution to the so-called ‘over-ionization problem’ in disc winds (REFs). This is the main motivation for incorporating microclumping into our model. This treatment is first-order; it does not adequately represent the complex substructures and stratifications in ionization state we expect in AGN outflows. Nevertheless, clumping is clearly important in these flows, and this parameterization allows a simple estimate for the effect clumping might have on the ionization state and emergent line emission. It is also encouraging that microclumping has been used successfully in fits to O-star wind spectra (Hillier 1991).

#### 4 RESULTS AND DISCUSSION

We need to describe what we did before getting to results. We started with parameters similar to those of H14, we created a small grid. Limits of the grid are given in table 1. This should be a paragraph or section before this one. Not all quasars have to have the same parameters and so we have to explain that we have elected to describe one model in detail. Was our goal to fix the luminosity and see what was required.

In next sections, we really need to deal with the difference between this class of models and “the model”, and especially “our model”. We may need to at least in words discuss whether we have “fine tuned” or not, by saying something about other models.

Here we describe the results from our model, the parameters of which are shown in table 1. Parameters differing from the benchmark model of H13 are highlighted with an asterisk. This set of parameters was arrived at by conducting a limited grid search over a 5-dimensional parameter space involving the variables  $r_{min}$ ,  $\theta_{min}$ ,  $f$ ,  $\alpha$  and  $R_V$ . We evaluated these models based on the following criteria:

- Does the model maintain the right ionization state to produce strong BALs ( $BI > ??$ )?
- Does significant line emission emerge at low inclinations, with  $EW \sim 50\text{\AA}$  in C IV?
- Do H recombination lines appear in the spectrum,  $EW \sim 50\text{\AA}$  in Ly  $\alpha$ ?
- Do a certain range of angles produce LoBAL features?
- Does the model compare favourably to quasar composite spectra?

In this section, we present one of the most promising models and discuss the various successes and failures. This allows us to gain insight into fundamental geometrical and physical constraints and assess the potential for unification. The full grid, including output synthetic spectra and plots can be found at [jhmatthews.github.io/quasar-wind-grid/](https://github.com/jhmatthews/quasar-wind-grid/).

##### 4.1 Physical Conditions and Ionization State

Figure 2 shows the physical properties of the wind. The wind rises slowly from the disc at first, with clumped densities of

Free Parameters	Value
$\dot{M}_{BH}$	$1 \times 10^9 M_\odot$
$\dot{M}_{acc}$	$5 M_\odot \text{yr}^{-1} \simeq 0.2 \dot{M}_{Edd}$
$\alpha_X$	$-0.9$
$L_X$	$10^{45} \text{ ergs s}^{-1}$ *
$r_{disc}(min) = r_X$	$6r_g = 8.8 \times 10^{14} \text{ cm}$
$r_{disc}(max)$	$3400r_g = 5 \times 10^{17} \text{ cm}$
$\dot{M}_{wind}$	$5 M_\odot \text{yr}^{-1}$
$r_{min}$	$300r_g = 4.4 \times 10^{16} \text{ cm}$
$r_{max}$	$600r_g = 8.8 \times 10^{16} \text{ cm}$
$\theta_{min}$	$70.0^\circ$
$\theta_{max}$	$82.0^\circ$
$\lambda$	$0$
$v_\infty(r_0)$	$v_{esc}(r_0)$
$R_v$	$10^{19} \text{ cm}$ *
$\alpha$	$0.6$ *
$f$	$0.01$ *

Table 1. Wind geometry parameters used in the model.

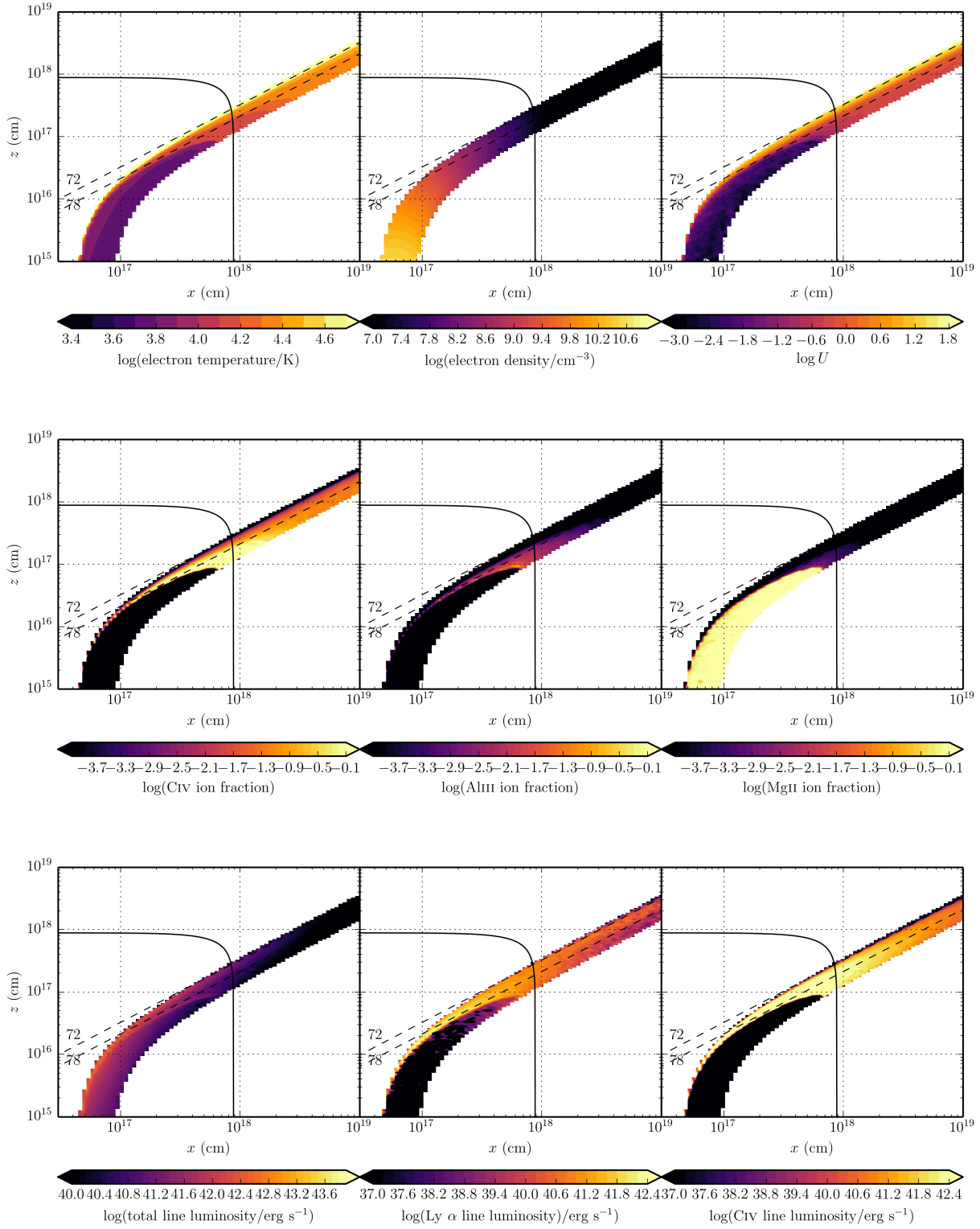
$n_H \sim 10^{11} \text{ cm}^{-3}$  close to the disc plane. The flow then accelerates over a scale length of  $R_V = 10^{19} \text{ cm}$  up to a terminal velocity equal to the escape velocity at the streamline base ( $\sim 10,000 \text{ km s}^{-1}$ ). This gradual acceleration means that the wind exhibits a stratified ionization structure, with low ionization material in the base of the wind giving way to highly ionized plasma further out. By clumping the wind, we are able to produce the range of ionization states observed in quasars and BALQSOs, while adopting a realistic  $2 - 10 \text{ keV}$  X-ray luminosity of  $L_X = 10^{45} \text{ ergs s}^{-1}$ . Without clumping, this wind would be over-ionized to the extent that opacities in e.g., C IV would be entirely negligible (see H13).

One common way to quantify the ionization state of a plasma is through the ionization parameter,  $U_H$ , given by

$$U_H = \frac{4\pi}{n_H c} \int_{13.6\text{eV}}^{\infty} \frac{J_\nu d\nu}{h\nu}. \quad (10)$$

where  $n_H$  is the local number density of H, and  $\nu$  denotes photon frequency. Shown in figure 2, the ionization parameter is a useful measure of the global ionization state, as it represents the ratio of the number density of H ionizing photons to the local H density. It is, however, a poor representation of the ionization state of species such as C IV as it encodes no information about the shape of the SED. In our case, the X-ray photons are dominant in the photoionization of the UV resonance line ions. This explains why a factor of 100 increase in X-ray luminosity requires a clumping factor of 0.01, even though the value of  $U_H$  decreases by only a factor of  $\sim 10$  compared to H13. This is also the reason for significant Lyman edge photoabsorption at the highest inclinations (see section 4.3).

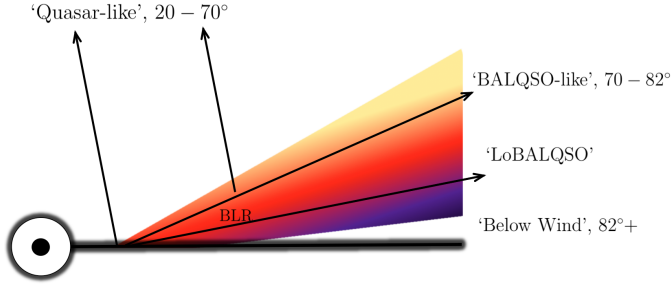
Clumping also causes the total line luminosity to increase dramatically, as recombination and collisional excitation are both proportional to  $n_e^2$ . This line emission typically emerges on the edge of the wind nearest the central source. The location of the line emitting regions is dependent on the ionization state, as well as the X-rays heating the plasma. The radii of these emitting regions is important, and can be compared to observations. As shown in figure 2, the C IV line in our model is typically formed be-



**Figure 2.** Physical properties of the outflow, shown by the coloured contours. The solid black line marks a sphere at  $1000 r_G$ . The dotted lines show the  $72^\circ$  and  $78^\circ$  sightlines to the centre of the system, and illustrate that different sightlines intersect material of different ionization states.

tween  $100 - 1000 r_G$  ( $\sim 10^{17} - 10^{18}$  cm). This is in rough agreement with the reverberation mapping results of Kaspi (2000) for the  $2.6 \times 10^9 M_\odot$  quasar S5 0836+71, and also compares favourably with microlensing measurements of the size of the C IV emission line region in the BALQSO H1413+117 (O’Dowd et al. 2015).

## 4.2 Synthetic Spectra: Comparison to Observations



**Figure 4.** A cartoon describing the broad classes of sightline in our model, illustrating how geometric effects lead to the different emergent spectra. The colour gradient is approximate, but indicates the stratified ionization structure, from highly ionized (yellow) to low ionization (purple) material.

Figure 3 shows the synthetic spectrum in the UV from our model. To assess the ability of the model to match real quasar spectra, we also show *Sloan Digital Sky Survey* (SDSS) quasar and BALQSO composites from Reichard et al. (2003), normalised to the flux at  $2000\text{\AA}$  in each panel. We show a cartoon illustrating how geometric effects determine the output spectra in figure 4.

### 4.2.1 Broad absorption lines

The UV spectrum is characterised by strong BAL profiles at high inclinations ( $> 70^\circ$ ). This highlights the first success of our model: clumping means the correct ionization state is maintained in the presence of strong X-rays, allowing large resonance line opacities. At the highest inclinations, the cooler, lower ionization material at the base of the wind starts to intersect the line of sight. This produces multiple absorption lines in lower ionization species such as Mg II, Al III and Fe II. The potential links to LoBALQSOs and FeLoBALQSOs are discussed in section 2.4.

The high ionization BAL profiles are often saturated, and the location in velocity space of the strongest absorption in the profile varies with inclination. At lower inclinations, the strongest absorption occurs at the red edge, whereas at high inclinations (and for the strongest BALs) the trough has a sharp edge at the terminal velocity. This offers one potential explanation for the wide range of BALQSO absorption line shapes (see e.g. Trump et al. 2006; Knigge et al. 2008). In addition, the line profile shape is strongly dependent on the density, ionization and velocity profiles intersected by the line of sight. Thus, small tweaks of the velocity law and angular distributions of streamlines can dramatically alter the shape of the line.

Nonblack saturation is observed in the absorption

troughs of BALQSOs (Arav et al. 1999b,a). This can be caused either by partial covering of the continuum source or by scattered contributions to the BAL troughs, necessarily from an opacity source not cospatial with the BAL forming region. The scattered light explanation is supported by spectropolarimetry results (Lamy & Hutsemékers 2000). Our spectra do not show nonblack saturation. Instead, we find black, saturated troughs at angles  $i > 73^\circ$ , and the BALs are non-saturated at lower inclinations. The reasons for this are readily apparent. First, the microclumping assumption does not allow for porosity in the wind, meaning that it does not naturally produce a partial covering absorber. To do this, an alternative approach such as *macro-clumping* would be required (e.g. Šurlan et al. 2012; Hamann et al. 2008). Second, our wind does not have a significant scattering contribution along sightlines which do not pass through the BAL region, meaning that any scattered contribution is obscured by the saturated troughs. This suggests that either the scattering cross-section of the wind must be increased (with higher mass loss rates or covering factors), or that an additional source of electron opacity is required, potentially in a polar direction above the disc.

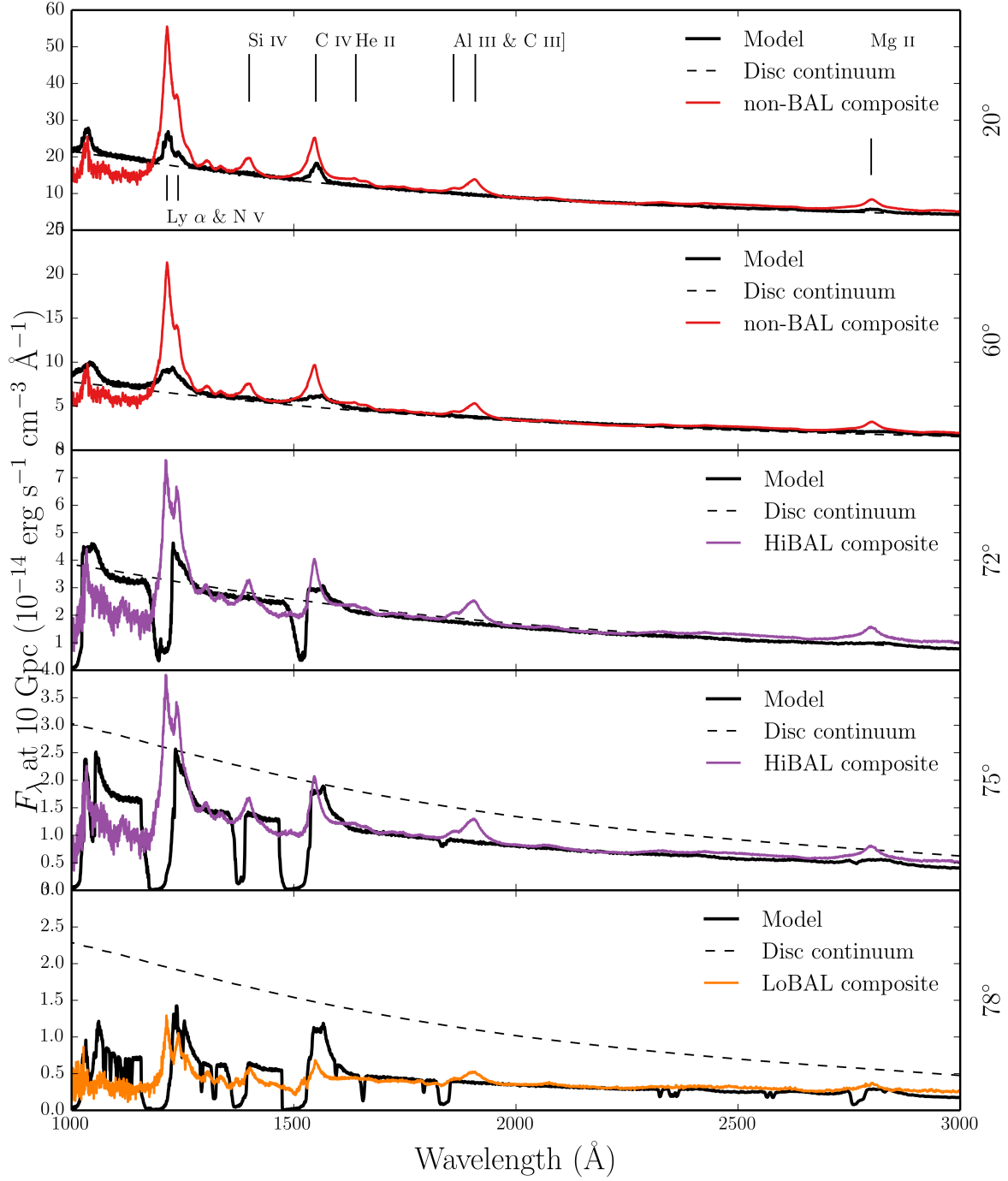
### 4.2.2 Broad emission lines

We find that the model can produce significant line emission at low inclinations, particular in C IV, and the improved treatment of recombination results in a strong Ly  $\alpha$  line. In the context of unification, this is a promising result, and shows that a biconical wind can produce significant emission at ‘quasar-like’ angles. Our simulated spectra do not contain the strong C III]  $1909\text{\AA}$  line seen in the quasar composite spectra. This is because we do not yet treat C as a full macro-atom with a full collisional rates between forbidden or semi-forbidden transitions, as would be required. The critical density of the C III]  $1909\text{\AA}$  line is  $n_e \sim 10^{9.5} \text{ cm}^{-3}$ , which is higher than much of the outer portion of our wind. We therefore expect a model with these parameters to produce C III]  $1909\text{\AA}$  with a proper treatment.

The model produces strong emission lines in C IV, N V and Ly  $\alpha$ , as well as a weak Mg II line. The shapes and widths of these lines match the composites fairly well. However, the line-to-continuum ratios at low inclinations in our model are significantly weaker than the quasar composites. Increasing the density of the outflow, by altering the mass loss rate or velocity law, can produce more line emission. However, the red wing of the BAL profiles is generally stronger than seen in BALQSO spectra and composites. This illustrates a fundamental problem with a geometric unification model such as this: that the line-to-continuum ratios at low inclinations are significantly affected by disc foreshortening and limb darkening. The angular distribution of the disc radiation is clearly crucially important in determining the emergent line ratios.

### 4.2.3 The Angular Distribution of Line And Continuum Emission

**: I am not convinced this is in the right place in the paper, because here we explore an alternative hypothesis. It might go better at the end. But a good many of the following sections are not that well motivated.**



**Figure 3.** Synthetic spectra at five viewing angles in our model. The coloured lines show different quasar and BAL quasar composites, and the dotted line shows a disc only continuum to show the effect of the outflow on the continuum level.



When comparing BALQSO and quasar composites, it is apparent that they possess remarkably similar line strengths and widths (e.g. Weymann et al. 1991; Reichard et al. 2003). **I think we mean EW here; otherwise I don't understand what we are talking about. We could say "line shapes and equivalent widths". I also think this requires more than one sentence of introduction.** This presents a challenge to our model, as well as the geometric unification picture in general, because these quantities should vary with inclination angle. Limb darkening and foreshortening causes the disc continuum emission to be strongly anisotropic, yet the line emission in our models is much more isotropic. This has the effect of enhancing the line-to-continuum ratios at high inclinations. To construct a scenario where emission line equivalent widths are comparable at all inclinations requires significant fine-tuning when considering a classical thin disc. **Is the previous sentence a throwaway line, or has some person done this. If so, please reference; if not please delete.**

Obviously, one possible solution to this problem is hypothesize a more isotropic distribution for the emergent condition than predicted by a "classical" disc. General relativistic effects – specifically, light bending and relativistic beaming – can cause the accretion disc SED to become more isotropic (e.g. Zhang et al. 1997; Muñoz-Darias et al. 2013). Incorporating these effects into PYTHON would be a very difficult and time-consuming challenge. Therefore, for the purposes of this discussion we have chosen to investigate whether GR effects could produce enough isotropy to address the EW problem, we have generated GR disc spectra, we use the code AGNSPEC (Hubeny et al. 2000; Davis & Hubeny 2006; Davis et al. 2007). The output flux at three different wavelengths as a function of inclination for an AGN-SPEC model with the same disc and BH parameters as our clumpy wind model is shown in figure ?? . The results indicate that the effects of GR on an AGN disc are much less extreme in the UV portion of the spectrum than the calculations by Zhang et al. (1997) for the X-rays in X-ray binaries. The emission for the disks of AGN is still strongly anisotropic. GR alone therefore cannot explain the line ratio trends in quasars by making the disc emit more isotropically. This supports the findings of Risaliti et al. (2011), who find that EW distributions in quasars are consistent with anisotropic emission from optically thick, disc-like structures for *both* the continuum source and BLR. If this is the case, it has a dramatic affect on the intrinsic BAL fraction inferred from flux-limited samples (Goodrich 1997; Krolik & Voit 1998). We will explore these ideas further in a future study. **Wait! All this says is that there has to be an alternative explanation. It is not support for any specific explanation!. This last few sentences should be a separate paragraph. They should start out: An alternative explanation proposed by risalite is that...**

### 4.3 X-ray Properties and Broadband SEDs

**: Again this section needs an introduction. Why is this important?**

Figure 5 shows the emergent monochromatic luminosity ( $L_\nu$ ) at 2 keV and plotted against  $L_\nu$  at 2500Å for a number of different viewing angles in our model. The monochromatic luminosities are calculated from the synthetic spectra

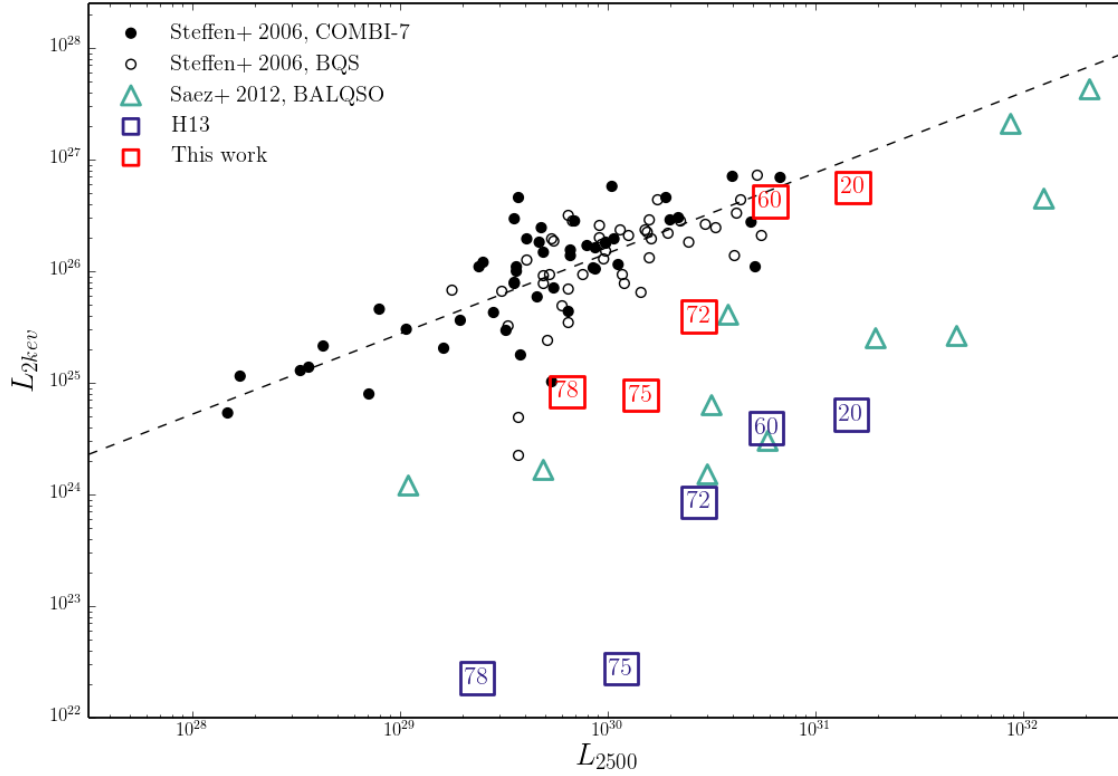
and thus include the effects of wind reprocessing and attenuation. In addition to model outputs, we also show the BALQSO sample of Saez et al. (2012) and luminous AGN and quasar samples from Steffen et al. (2006). The best fit relation from Steffen et al. (2006) is also shown. For low inclination, 'quasar-like' viewing angles, we now show excellent agreement with AGN samples. The gradient from  $20^\circ$  to  $60^\circ$  in our models is caused by a combination of disc foreshortening/limb-darkening (resulting in a lower  $L_{2500}$  for higher inclinations) and the fact that the disk is opaque, and thus the X-ray source subtends a smaller solid angle at high inclinations (resulting in a lower  $L_{2keV}$  for higher inclinations).

The low inclination, 'BALQSO-like' viewing angles show moderate agreement with the data, and are X-ray weak due to bound-free and electron scattering opacities in the wind. Typically, BALQSOs show strong X-ray absorption with columns of  $N_H \sim 10^{23} \text{ cm}^{-2}$  (Green & Mathur 1996; Mathur et al. 2000; Green et al. 2001; Grupe et al. 2003). This is often cited as evidence that the BAL outflow is shielded from the X-ray source, especially as sources with strong X-ray absorption tend to exhibit deep BAL troughs and high outflow velocities (Brandt et al. 2000; Laor & Brandt 2002; Gallagher et al. 2006). Our results imply that the clumpy BAL outflow itself can be responsible for the strong X-ray absorption, and supports Hamann et al.'s (2013) suggestion that this explains the weaker X-ray absorption in mini-BALs compared to BALQSOs.

Our models slightly over-predict the emergent X-ray luminosity at BAL angles, although we are limited by poor sample sizes. If BALQSOs were *intrinsically* X-ray weak (as suggested by, e.g. Morabito et al. 2013), our isotropic assumption for the X-ray source would be incorrect. A polar-biased X-ray source would result in a lower clumping factor being required in our model. Our specific wind prescription will also affect the opacities, densities and resultant ionization structure, which can change the absorption characteristics and resultant luminosities. Nevertheless, our input X-ray spectrum now reproduces the X-ray properties of a luminous quasar as an output, and at least some BAL angles match the observations. This satisfies the first-order requirement for the X-ray properties of a unified quasar model.

### 4.4 LoBALs and ionization stratification

At certain sightlines, our model now produces blue-shifted BALs in Al III and Mg II – the absorption lines seen in LoBALQSOs, and we even see absorption in Fe II at the highest inclinations. Line profiles in velocity space for C IV, Al III and Mg II, are shown in figure 6 for a range of BALQSO viewing angles. We find that ionization stratification of the wind causes lower ionization material has a smaller covering factor, as demonstrated by figures 2 and 6. This confirms the behaviour expected from a unification model such as Elvis (2000). LoBALs are only present at viewing angles close to edge-on ( $i > 75^\circ$ ), as predicted by polarisation results (Brotherton et al. 1997). There is also a correlation between the strength of LoBAL features and the amount of continuum attenuation at that sightline, particularly blueward of the Lyman edge as the low ionization base intersects the line-of-sight. Our model therefore predicts that LoBALQSOs and FeLoBALQSOs have stronger Lyman edge



**Figure 5.** X-ray (2 keV) luminosity of the our clumped model (red squares) and the H13 model (purple squares), plotted against monochromatic luminosity at 2500Å. The points are labeled according to inclination; angles  $> 70^\circ$  correspond to BALs in our scheme (see figure 4). Also plotted are the samples considered by Saez et al. 2012 on a similar plot; The COMBI-7 AGN and the BQS samples Steffen et al. (2006) and the Saez et al. (2012) sample of BALQSOs. The dotted line shows the best fit relation for non-BALQSOs from Steffen et al. (2006).

absorption and be more Compton-thick than HiBALQSOs and Type 1 quasars. An edge-on scenario also offers a potential explanation for the rarity of LoBAL and FeLoBAL quasars, due to a foreshortened and attenuated continuum, although, as noted in section 4.2.3, BAL fraction inferences are fraught with complex selection effects.

**Note: One of the things I definitely wondered about is whether our "model" would produce the fraction of BALs seen in the sky**

## 5 SUMMARY

We have carried out MCRT simulations using a simple prescription for a biconical disc wind, with the aim of expanding on the work of H13 and assessing the viability of such a model for geometric unification of quasars. We find the following main points:

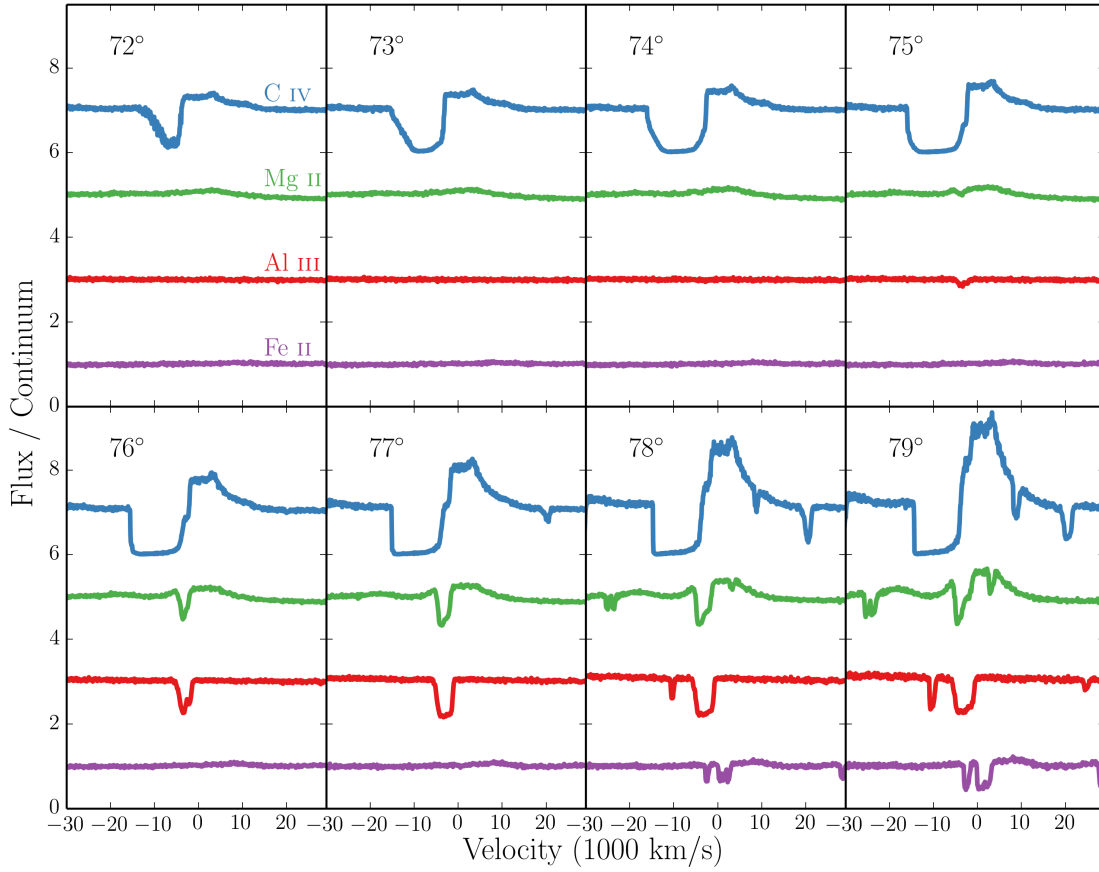
- (i) We have introduced a first-order treatment of clumping in our model, and found that it can now maintain the required ionization state while agreeing well with the X-ray properties of AGN/QSOs.
- (ii) We have shown that the degree of ionization stratification in the model is sufficient that LoBAL line profiles are

seen at a subset of viewing angles, and Fe II absorption is seen at particularly high inclinations.

- (iii) We find that clumping also causes a significant increase in the strength of the emission lines produced by the model. This is true both of collisionally excited resonance lines (such as C IV, N V) and recombination lines (such as Ly  $\alpha$ , H  $\alpha$  and the Balmer series).

- (iv) The line EWs in our models increase with inclination. BAL and non-BAL quasar composites have comparable EWs, so our model fails to reproduce this behaviour. This is due to a fundamental constraint discussed further in section 5. If the BLR emits fairly isotropically then for a foreshortened, limb-darkened classical thin accretion disc it is simply not possible to achieve line ratios at low inclinations that are comparable to those at high inclinations. This is a robust conclusion which is independent of the assumed BLR geometry and size.

- (v) We have examined the effect of GR on our disc SED, using the disc atmosphere and GR ray-tracing code AGNSPEC. While including GR effects does cause the disc SED to become slightly more isotropic, the effect is not large enough



**Figure 6.** C IV, Mg II, Al III and Fe II line profiles for wind angles from 72 – 79°. The profiles are plotted relative to the local continuum with an offset applied for clarity. Lower ionization profiles appear at a subset of high inclinations, compared to the ubiquitous C IV profile.

to produce uniform line to continuum ratios with viewing angle. We briefly discuss other solutions.

Our work confirms a number of expected outcomes from a geometric unification model, and suggests that a simple biconical geometry such as this can come close to explaining much of the phenomenology of quasars. Nevertheless, our conclusions pose a clear challenge to the current disc wind unification picture.

## ACKNOWLEDGEMENTS

The work of JHM, SWM, NSH and CK is supported by the Science and Technology Facilities Council (STFC), via two studentships and a consolidated grant, respectively. We would like to thank Omer Blaes, Ivan Hubeny and Shane Davis for their assistance with AGNSPEC. We would also like to thank Daniel Proga, Daniel Capellupo, Sam Connolly and Dirk Grupe for useful discussions. Simulations were conducted using PYTHON version 79c, and made use of the IRIDIS High Performance Computing Facility at the University of Southampton. Figures were produced using matplotlib (Hunter 2007).

## REFERENCES

- Allen J. T., Hewett P. C., Maddox N., Richards G. T., Belokurov V., 2011, *MNRAS* 410, 860
- Arav N., 1996, *ApJ* 465, 617
- Arav N., Becker R. H., Laurent-Muehleisen S. A., Gregg M. D., White R. L., Brotherton M. S., de Kool M., 1999a, *ApJ* 524, 566
- Arav N., Korista K. T., Barlow T. A., Begelman, 1995, *Nature* 376, 576
- Arav N., Korista K. T., de Kool M., Junkkarinen V. T., Begelman M. C., 1999b, *ApJ* 516, 27
- Badnell N. R., 2006, *ApJs* 167, 334
- Begelman M., de Kool M., Sikora M., 1991, *ApJ* 382, 416
- Blaes O., 1998, in S. S. Holt, T. R. Kallman (eds.), *American Institute of Physics Conference Series*, Vol. 431 of *American Institute of Physics Conference Series*, p. 161
- Blandford R. D., Payne D. G., 1982, *MNRAS* 199, 883
- Bowler R. A. A., Hewett P. C., Allen J. T., Ferland G. J., 2014, *MNRAS* 445, 359
- Brandt W. N., Laor A., Wills B. J., 2000, *ApJ* 528, 637
- Brotherton M. S., Tran H. D., van Breugel W., Dey A.,

- Antonucci R., 1997, *ApJ Letters* 487, L113
- Capellupo D. M., Hamann F., Barlow T. A., 2014, *MNRAS* 444, 1893
- Capellupo D. M., Hamann F., Shields J. C., Rodríguez Hidalgo P., Barlow T. A., 2011, *MNRAS* 413, 908
- Capellupo D. M., Hamann F., Shields J. C., Rodríguez Hidalgo P., Barlow T. A., 2012, *MNRAS* 422, 3249
- Capellupo D. M., Netzer H., Lira P., Trakhtenbrot B., Mejía-Restrepo J., 2015, *MNRAS* 446, 3427
- Carlberg R. G., 1980, *ApJ* 241, 1131
- Cassidy I., Raine D. J., 1996, *A&A* 310, 49
- Cohen M. H., Ogle P. M., Tran H. D., Vermeulen R. C., Miller J. S., Goodrich R. W., Martel A. R., 1995, *ApJ Letters* 448, L77
- Cottis C. E., Goad M. R., Knigge C., Scaringi S., 2010, *MNRAS* 406, 2094
- Cunto W., Mendoza C., Ochsenbein F., Zeippen C. J., 1993, *A&A* 275, L5
- Davis S. W., Hubeny I., 2006, *ApJs* 164, 530
- Davis S. W., Woo J.-H., Blaes O. M., 2007, *ApJ* 668, 682
- de Kool M., Begelman M. C., 1995, *ApJ* 455, 448
- Dere K. P., 2007, *A&A* 466, 771
- Dere K. P., Landi E., Mason H. E., Monsignori Fossi B. C., Young P. R., 1997, *A&As* 125, 149
- Elitzur M., Ho L. C., Trump J. R., 2014, *MNRAS* 438, 3340
- Elvis M., 2000, *ApJ* 545, 63
- Emmering R. T., Blandford R. D., Shlosman I., 1992, *ApJ* 385, 460
- Fabian A. C., 2012, *ARAA* 50, 455
- Gallagher S. C., Brandt W. N., Chartas G., Priddey R., Garmire G. P., Sambruna R. M., 2006, *ApJ* 644, 709
- Ganguly R., Brotherton M. S., 2008, *ApJ* 672, 102
- Ganguly R., Sembach K. R., Tripp T. M., Savage B. D., Wakker B. P., 2006, *ApJ* 645, 868
- Goodrich R. W., 1997, *ApJ* 474, 606
- Goodrich R. W., Miller J. S., 1995, *ApJ Letters* 448, L73
- Green P. J., Aldcroft T. L., Mathur S., Wilkes B. J., Elvis M., 2001, *ApJ* 558, 109
- Green P. J., Mathur S., 1996, *ApJ* 462, 637
- Grupe D., Mathur S., Elvis M., 2003, *AJ* 126, 1159
- Hamann F., Chartas G., McGraw S., Rodríguez Hidalgo P., Shields J., Capellupo D., Charlton J., Eracleous M., 2013, *MNRAS* 435, 133
- Hamann W.-R., Koesterke L., 1998, *A&A* 335, 1003
- Hamann W.-R., Oskinova L. M., Feldmeier A., 2008, in W.-R. Hamann, A. Feldmeier, L. M. Oskinova (eds.), *Clumping in Hot-Star Winds*, 75
- Häring N., Rix H.-W., 2004, *ApJ Letters* 604, L89
- Hazard C., Mackey M. B., Shimmings A. J., 1963, *Nature* 197, 1037
- Higginbottom N., Knigge C., Long K. S., Sim S. A., Matthews J. H., 2013, *MNRAS* 436, 1390
- Higginbottom N., Proga D., Knigge C., Long K. S., Matthews J. H., Sim S. A., 2014, *ApJ* 789, 19
- Hillier D. J., 1991, *A&A* 247, 455
- Hubeny I., Agol E., Blaes O., Krolik J. H., 2000, *ApJ* 533, 710
- Hunter J. D., 2007, *Computing In Science & Engineering* 9(3), 90
- Kellermann K. I., Sramek R., Schmidt M., Shaffer D. B., Green R., 1989, *AJ* 98, 1195
- King A., 2003, *ApJ Letters* 596, L27
- King A., 2005, *ApJ Letters* 635, L121
- Knigge C., 1999, *MNRAS* 309, 409
- Knigge C., Scaringi S., Goad M. R., Cottis C. E., 2008, *MNRAS* 386, 1426
- Krolik J. H., McKee C. F., Tarter C. B., 1981, *ApJ* 249, 422
- Krolik J. H., Voit G. M., 1998, *ApJ Letters* 497, L5
- Lamy H., Hutsemékers D., 2000, *A&A* 356, L9
- Landi E., Del Zanna G., Young P. R., Dere K. P., Mason H. E., 2012, *ApJ* 744, 99
- Laor A., Brandt W. N., 2002, *ApJ* 569, 641
- Laor A., Davis S. W., 2014, *MNRAS* 438, 3024
- Long K. S., Knigge C., 2002, *ApJ* 579, 725
- Lucy L. B., 2002, *A&A* 384, 725
- Lucy L. B., 2003, *A&A* 403, 261
- Lucy L. B., Solomon P. M., 1970, *ApJ* 159, 879
- MacGregor K. B., Hartmann L., Raymond J. C., 1979, *ApJ* 231, 514
- Marscher A. P., 2006, in P. A. Hughes, J. N. Bregman (eds.), *Relativistic Jets: The Common Physics of AGN, Microquasars, and Gamma-Ray Bursts*, Vol. 856 of *American Institute of Physics Conference Series*, p. 1
- Mathur S., Green P. J., Arav N., Brotherton M., Crenshaw M., deKool M., Elvis M., Goodrich R. W., Hamann F., Hines D. C., Kashyap V., Korista K., Peterson B. M., Shields J. C., Shlosman I., van Breugel W., Voit M., 2000, *ApJ Letters* 533, L79
- Matthews J. H., Knigge C., Long K. S., Sim S. A., Higginbottom N., 2015, *MNRAS* 450, 3331
- Morabito L. K., Dai X., Leighly K. M., Sivakoff G. R., Shankar F., 2013, *ArXiv e-prints*
- Muñoz-Darias T., Coriat M., Plant D. S., Ponti G., Fender R. P., Dunn R. J. H., 2013, *MNRAS* 432, 1330
- Murray N., Chiang J., Grossman S. A., Voit G. M., 1995, *ApJ* 451, 498
- Noebauer U. M., Long K. S., Sim S. A., Knigge C., 2010, *ApJ* 719, 1932
- North M., Knigge C., Goad M., 2006, *MNRAS* 365, 1057
- O'Dowd M. J., Bate N. F., Webster R. L., Labrie K., Rogers J., 2015, *ArXiv e-prints*
- Owocki S. P., Rybicki G. B., 1984, *ApJ* 284, 337
- Owocki S. P., Rybicki G. B., 1985, *ApJ* 299, 265
- Pelletier G., Pudritz R. E., 1992, *ApJ* 394, 117
- Perley R. A., Dreher J. W., Cowan J. J., 1984, *ApJ Letters* 285, L35
- Potash R. I., Wardle J. F. C., 1980, *ApJ* 239, 42
- Pounds K. A., Reeves J. N., 2009, *MNRAS* 397, 249
- Proga D., Jiang Y.-F., Davis S. W., Stone J. M., Smith D., 2014, *ApJ* 780, 51
- Proga D., Kallman T. R., 2004, *ApJ* 616, 688
- Proga D., Kurosawa R., 2010, in L. Maraschi, G. Ghisellini, R. Della Ceca, F. Tavecchio (eds.), *Accretion and Ejection in AGN: a Global View*, Vol. 427 of *Astronomical Society of the Pacific Conference Series*, 41
- Proga D., Stone J. M., Kallman T. R., 2000, *ApJ* 543, 686
- Reeves J. N., O'Brien P. T., Ward M. J., 2003, *ApJ Letters* 593, L65
- Reichard T. A., Richards G. T., Hall P. B., Schneider D. P., Vanden Berk D. E., Fan X., York D. G., Knapp G. R., Brinkmann J., 2003, *AJ* 126, 2594
- Risaliti G., Elvis M., Nicastro F., 2002, *ApJ* 571, 234

- Risaliti G., Salvati M., Marconi A., 2011, MNRAS 411, 2223
- Shakura N. I., Sunyaev R. A., 1973, A&A 24, 337
- Shlosman I., Vitello P., 1993, ApJ 409, 372
- Shlosman I., Vitello P. A., Shaviv G., 1985, ApJ 294, 96
- Silk J., Rees M. J., 1998, A&A 331, L1
- Sim S. A., Drew J. E., Long K. S., 2005, MNRAS 363, 615
- Sim S. A., Long K. S., Miller L., Turner T. J., 2008, MNRAS 388, 611
- Sim S. A., Miller L., Long K. S., Turner T. J., Reeves J. N., 2010, MNRAS 404, 1369
- Simon L. E., Hamann F., 2010, MNRAS 409, 269
- Springel V., Di Matteo T., Hernquist L., 2005, ApJ Letters 620, L79
- Sutherland R. S., 1998, MNRAS 300, 321
- Tombesi F., Cappi M., Reeves J. N., Palumbo G. G. C., Yaqoob T., Braitto V., Dadina M., 2010, A&A 521, A57
- Turner T. J., Miller L., 2009, AAPR 17, 47
- Šurlan B., Hamann W.-R., Kubát J., Oskinova L. M., Feldmeier A., 2012, A&A 541, A37
- Verner D. A., Ferland G. J., Korista K. T., Yakovlev D. G., 1996, ApJ 465, 487
- Weymann R. J., Morris S. L., Foltz C. B., Hewett P. C., 1991, ApJ 373, 23
- Weymann R. J., Scott J. S., Schiano A. V. R., Christiansen W. A., 1982, ApJ 262, 497
- Zhang S. N., Cui W., Chen W., 1997, ApJ Letters 482, L155