

**Table 4**

<b>Molecular Hydrogen Transitions</b>	<b>Photon Energies (cm<sup>-1</sup>)</b>	<b>Relative Intensities</b>	<b>Indices <i>i</i></b>
(2,1)S(7)	5397.2	0.12	1
(9,7)S(3)	5324.8	0.05	2
(1,0)S(4)	5285.6	0.37	3
(2,1)S(6)	5277.8	0.08	4
(9,7)S(1)	5146.8	0.11	5
(2,1)S(5)	5141.8	0.25	6
(1,0)S(3)	5108.4	1.07	7
(9,7)S(0)	5031.8	0.06	8
(2,1)S(4)	4989.8	0.19	9
(1,0)S(2)	4917.0	0.80	10
(3,2)S(5)	4841.3	0.11	11
(2,1)S(3)	4822.8	0.56	12
(1,0)S(1)	4712.9	1.60	13
(3,2)S(4)	4699.3	0.09	14
(2,1)S(2)	4642.1	0.44	15
(3,2)S(3)	4542.6	0.28	16
(1,0)S(0)	4497.8	0.73	17
(2,1)S(1)	4449.0	0.89	18
(3,2)S(2)	4372.4	0.23	19
(4,3)S(3)	4265.2	0.13	20

Molecular hydrogen (H<sub>2</sub>) transitions produce photons that cool Uranus's protosatellite disk at concentric rings. The radius of each ring depends on a particular photon energy in the H<sub>2</sub> spectrum.

In this table the *Molecular Hydrogen Transitions* are the names of the transitions that produce the photons with *photon energies* as listed. These photons are the ones that initiated the evolution of the satellites of Uranus, Jupiter and Neptune. It is not necessary to understand the nomenclature for these transition names.

For simplicity the indices (*i*) have been introduced and each one refers to a particular photon energy in the H<sub>2</sub> spectrum.

For a discussion of *relative intensities*, see the text accompanying this table.