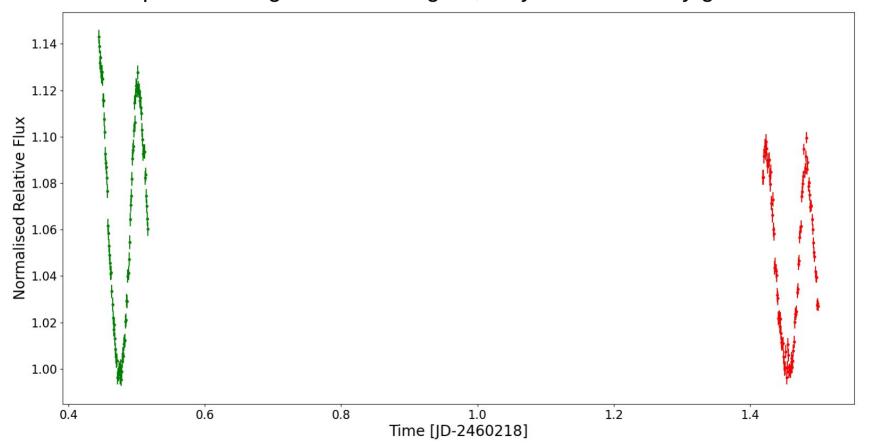
The variable star data were taken on two different nights, so when we plot them together in one figure, they don't look very good...



#### **Ephemeris**

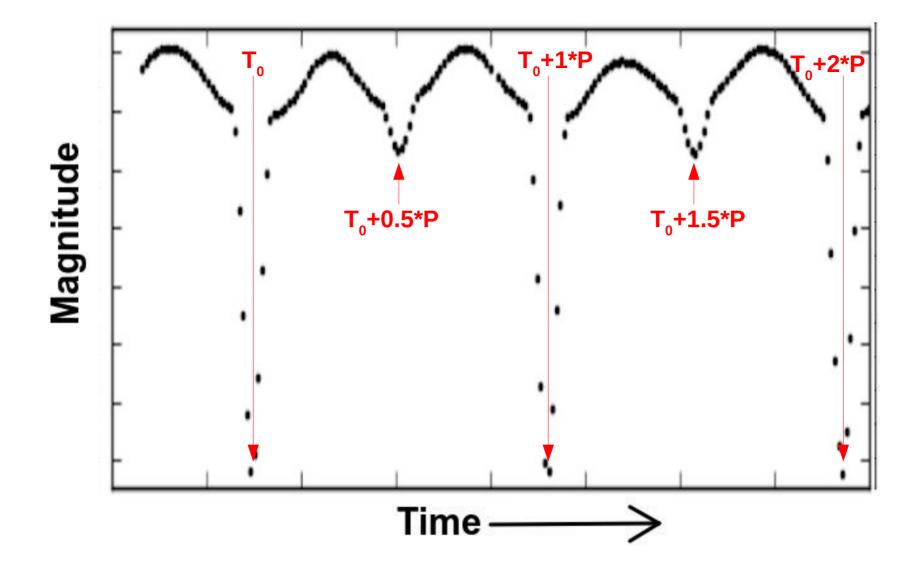
$$T_{obs} = T_0 + E*P$$

- > T<sub>obs</sub> = the timestamp of each data point, usually in some JD format (JD, MJD, HJD, BMJD...)
- $T_0$  = a reference time, e.g. the first observed time of mid-eclipse the first observed pulsation maximum
- E = the cycle number (just that, a number)
- P = the variability period (pulsation period, orbital period etc)

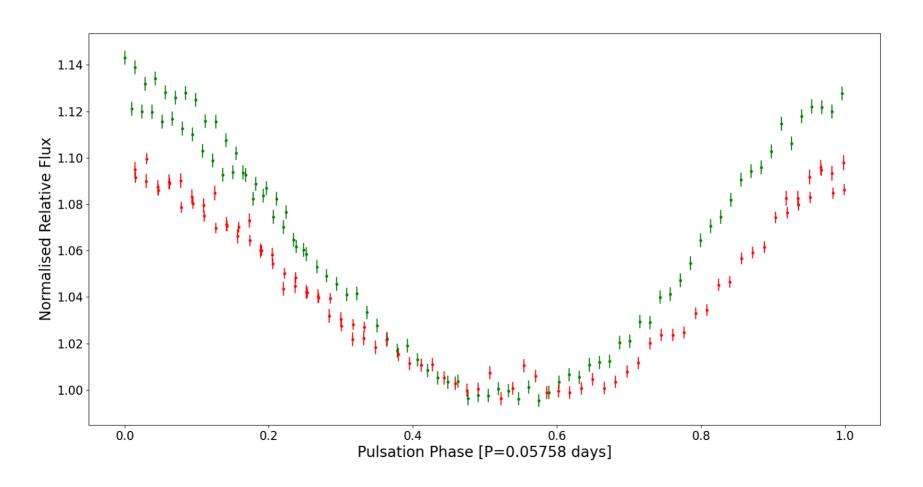
## Phase-folding data

$$E = \frac{T_{obs} - T_0}{P}$$

- $\succ$  Calculating E allows us to see how many full periods have passed between T<sub>0</sub> and T<sub>obs</sub> (the integer part of the resulting E)
- It also allows us to calculate on what phase (= percentage of one full cycle) our observations were taken (the decimal part)



#### Pulsation Period: 0.0575814 days



## Getting the magnitude

$$mag_{T} = \frac{-\ln(\sum_{i} 2.512^{-mag_{C_{i}}})}{\ln(2.512)} - 2.5\log(rel_{flux_{T}})$$

Where mag<sub> $c_i$ </sub> is the catalogue magnitude of the i-th comparison star (i=2, 3, 4...)

rel\_flux\_T you can get directly from your full Table of measurements and it is defined as:

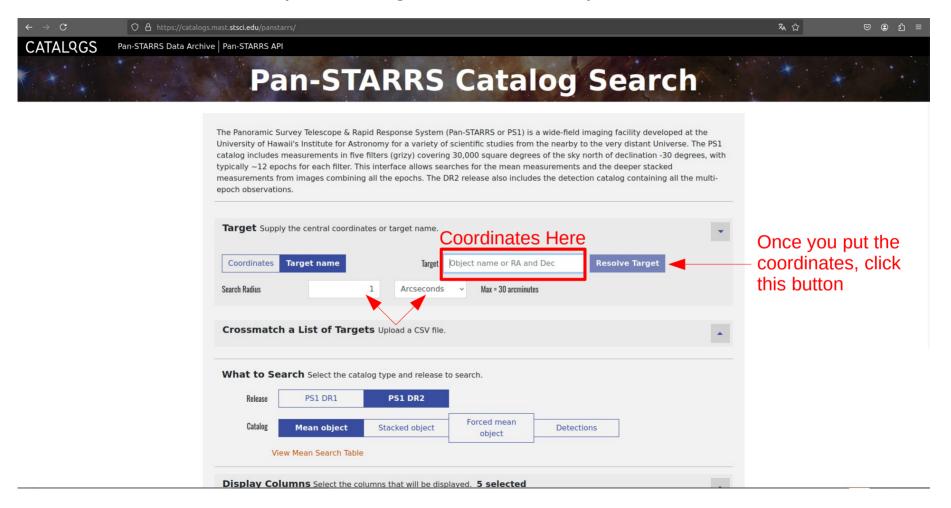
$$rel_flux_T = \frac{Source - Sky_T 1}{\sum_{i} Souce - Sky_C_i}$$

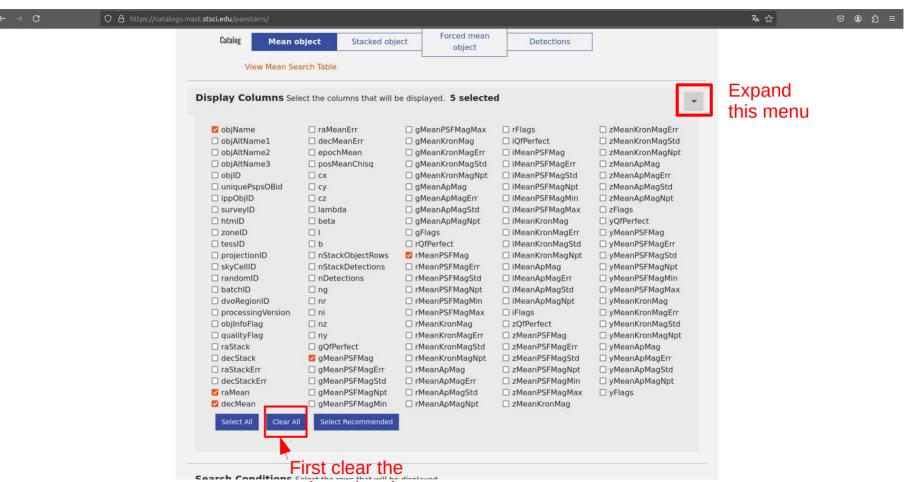
# Getting the magnitude error

$$Mag\_Err_T = 2.5*log \begin{vmatrix} 1 + \sqrt{\frac{Source\_Error\_T1^2}{Source-Sky\_T1^2}} + \frac{\sum_{i} Source\_Err\_C_i^2}{\left|\sum_{i} Source-Sky\_C_i\right|^2} \end{vmatrix}$$

All the different variables that enter in the Equation above, you can get directly from the AstroImageJ Measurements Table!

#### https://catalogs.mast.stsci.edu/panstarrs/





selected options