

MATHEMATICS

Trig

Identities:

sin(A ± B) = sin A cos B ± cos A sin B
cos(A ± B) = cos A cos B ∓ sin A sin B

sin A sin B = 1/2 [cos(A - B) - cos(A + B)]
cos A cos B = 1/2 [cos(A - B) + cos(A + B)]
sin A cos B = 1/2 [sin(A + B) + sin(A - B)]
cos A sin B = 1/2 [sin(A + B) - sin(A - B)]

sin^2 A = (1 - cos 2A) / 2
cos^2 A = (1 + cos 2A) / 2

sin A + sin B = 2 sin((A+B)/2) cos((A-B)/2)
sin A - sin B = 2 cos((A+B)/2) sin((A-B)/2)
cos A + cos B = 2 cos((A+B)/2) cos((A-B)/2)
cos A - cos B = -2 sin((A+B)/2) sin((A-B)/2)

Law of Sines: sides: A, B, and C; angles opposite: α, β, and γ

A/sin α = B/sin β = C/sin γ

Law of Cosines: sides: A, B, and C; angle opposite of C = γ

C^2 = A^2 + B^2 - 2AB cos γ

Circular Arclength: s = rθ; θ in rad

Vectors

Divergence theorem (Gauss' thm): ∫_V (∇ · F) dV = ∫_S F · n dS
Stokes theorem: ∫_S ∇ × F · dS = ∮_C F · dr
Dot product: A · B = |A||B| cos θn = A_x B_x + A_y B_y + A_z B_z
Cross product: A × B = |A||B| sin θn

Coordinate systems & conversions

COORDINATE CONVERSION		
cartesian to...	cylindrical to...	spherical to...
cylindrical	cartesian	cartesian
ρ = √(x² + y²) φ = arctan(y/x) z = z	x = ρ cos φ y = ρ sin φ z = z	x = r sin θ cos φ y = r sin θ sin φ z = r cos θ
spherical	spherical	cylindrical
r = √(x² + y² + z²) θ = arccos(z/r) φ = arctan(y/x)	r = √(ρ² + z²) θ = arctan(ρ/z) φ = φ	ρ = r sin θ φ = r sin θ z = r cos θ

UNIT VECTOR CONVERSION

cartesian to...	cylindrical to...	spherical to...
cylindrical	cartesian	cartesian
ρ̂ = x/ρ î + y/ρ ĵ	î = cos φ ρ̂ - sin φ φ̂	î = sin θ cos φ ρ̂ + cos θ cos φ θ̂ - sin φ φ̂
φ̂ = -y/ρ î + x/ρ ĵ	î = cos φ ρ̂ - sin φ φ̂	ĵ = sin θ sin φ ρ̂ + cos θ sin φ θ̂ + cos φ φ̂
ẑ = ẑ	ẑ = ẑ	ẑ = cos θ ρ̂ - sin θ θ̂
spherical	spherical	cylindrical
ρ̂ = (xî + yĵ + ẑ)/ρ	ρ̂ = ρ/ρ̂ ρ̂ + z/ρ̂ ẑ	ρ̂ = sin θ ρ̂ + cos θ θ̂
θ̂ = (xzî + yzĵ - ρ² ẑ)/ρ²	θ̂ = z/ρ̂ ρ̂ - ρ/ρ̂ ẑ	φ̂ = φ̂
φ̂ = (-yî + xĵ)/ρ	φ̂ = φ̂	ẑ = cos θ ρ̂ - sin θ θ̂

DIFFERENTIAL ELEMENTS		
cartesian	cylindrical	spherical
dî = dxî + dyĵ + dẑ	dî = dρ ρ̂ + ρ dφ φ̂ + dz ẑ	dî = dr ρ̂ + r dθ θ̂ + r sin θ dφ φ̂
dA = dy dz î + dx dz ĵ + dx dy ẑ	dA = ρ dy dz ρ̂ + dρ dz φ̂ + ρ dρ dφ ẑ	dA = r² sin θ dθ dφ ρ̂ + r sin θ dr dφ θ̂ + r dr dθ φ̂
dV = dx dy dz	dV = ρ dρ dφ dz	dV = r² sin θ dr dθ dφ

POS, VEL, ACCEL

polar:
r = ρe_ρ
v = ρ̇e_ρ + ρθ̇e_θ
a = (ρ̈ - ρθ̇²)e_ρ + (ρθ̈ + 2ρ̇θ̇)e_θ

cylindrical:
r = ρe_ρ + ze_z
v = ρ̇e_ρ + ρθ̇e_θ + że_z
a = (ρ̈ - ρθ̇²)e_ρ + (ρθ̈ + 2ρ̇θ̇)e_θ + z̈e_z

spherical:
r = ρe_ρ + ρθ̇e_θ + ρφ̇ sin θ e_φ
a = (ρ̈ - ρθ̇² - ρφ̇² sin² θ)e_ρ + (ρθ̈ + 2ρ̇θ̇ - ρφ̇² sin θ cos θ)e_θ + (ρφ̈ sin θ + 2ρ̇φ̇ sin θ + 2ρθ̇φ̇ cos θ)e_φ

del, ∇, in CARTESIAN

del operator: ∇ = e_x ∂/∂x + e_y ∂/∂y + e_z ∂/∂z
gradient: ∇φ = grad φ = e_x ∂φ/∂x + e_y ∂φ/∂y + e_z ∂φ/∂z
directional derivative: dφ/ds = ∇φ · A/|A|
divergence: ∇ · V = div V = ∂V_x/∂x + ∂V_y/∂y + ∂V_z/∂z
curl: ∇ × V = e_x (∂V_z/∂y - ∂V_y/∂z) + e_y (∂V_x/∂z - ∂V_z/∂x) + e_z (∂V_y/∂x - ∂V_x/∂y)
Laplacian: Δf = ∇²φ = ∇ · (∇φ) = div grad φ = ∂²φ/∂x² + ∂²φ/∂y² + ∂²φ/∂z²

del, ∇, in CYLINDRICAL

gradient: ∇f = grad f = ∂f/∂ρ e_ρ + 1/ρ ∂f/∂φ e_φ + ∂f/∂z e_z
divergence: ∇ · V = div V = 1/ρ ∂(ρV_ρ)/∂ρ + 1/ρ ∂V_φ/∂φ + ∂V_z/∂z
curl: ∇ × V = (1/ρ ∂V_z/∂φ - ∂V_φ/∂z) e_ρ + (∂V_ρ/∂z - ∂V_z/∂ρ) e_φ + 1/ρ (∂(ρV_φ)/∂ρ - ∂V_ρ/∂φ) e_z
Laplacian: Δf = ∇²f = ∇ · (∇f) = 1/ρ ∂/∂ρ (ρ ∂f/∂ρ) + 1/ρ² ∂²f/∂φ² + ∂²f/∂z²

del, ∇, in SPHERICAL

gradient: ∇f = grad f = ∂f/∂r e_r + 1/r ∂f/∂θ e_θ + 1/(r sin θ) ∂f/∂φ e_φ
divergence: ∇ · V = div V = 1/r² ∂(r²V_r)/∂r + 1/(r sin θ) ∂(V_θ sin θ)/∂θ + 1/(r sin θ) ∂V_φ/∂φ
curl: ∇ × V = 1/(r sin θ) (∂/∂θ (V_φ sin θ) - ∂V_θ/∂φ) e_r + 1/r (1/sin θ ∂V_r/∂φ - ∂/∂r (rV_φ)) e_θ + 1/r (∂/∂r (rV_θ) - ∂V_r/∂θ) e_φ
Laplacian: Δf = ∇²f = ∇ · (∇f) = 1/r² ∂/∂r (r² ∂f/∂r) + 1/(r² sin θ) ∂/∂θ (sin θ ∂f/∂θ) + 1/(r² sin² θ) ∂²f/∂φ²

Taylor series

Taylor series of f(x) about x = a:

f(x) = f(a) + (x - a)f'(a) + 1/2! (x - a)² f''(a) + ... + 1/n! (x - a)^n f^(n)(a) + ...

Ordinary differential equations

Separable 1st-order

Equation can be written as f(y)dy = f(x)dx, such as dy/dx = N(1 - y). Evaluate integrals directly.

Linear 1st-order

Write the equation in the form y' + P(x)y = Q(x) and then define I = ∫ P(x)dx and find y by solving ye^I = ∫ Q(x)e^I dx + c

Linear 2nd-order homogeneous with constant coefficients

Equations of the form a₂ d²y/dx² + a₁ dy/dx + a₀ y = 0
a and b:

a, b ∈ ℝ, a ≠ b ⇒ y = c₁e^{ax} + c₂e^{bx}

a, b ∈ ℝ, a = b ⇒ y = (Ax + B)e^{ax}

For a, b ∈ ℑ, a = b* = α ± iβ, any of the following forms are solutions

y = Ae^{α+iβx} + Be^{α-iβx}

y = e^{αx} (Ae^{iβx} + Be^{-iβx})

y = e^{αx} (c₁ sin βx + c₂ cos βx)

y = ce^{αx} sin(βx + γ)

y = ce^{αx} cos(βx + δ)

Linear 2nd-order inhomogeneous with constant coefficients

Equations of one of the forms

a₂ d²y/dx² + a₁ dy/dx + a₀ y = f(x)

d²y/dx² + a₁/dx dy + a₀/y = F(x)

(D - a)(D - b)y = F(x) = ke^{cx}, particular solution y_p is given by:

y_p = Ce^{cx} if c is not equal to either a or b;
y_p = Cxe^{cx} if c equals a or b, a ≠ b;
y_p = Cx²e^{cx} if c = a = b

(For F(x) = k cos ax or F(x) = k sin ax, solve the above with F(x) = ke^{c=iax} and take the real or imag part, respectively. For F(x) = const, set c = 0.)

A more general form of this (called the method of undetermined coefficients) follows:

(D - a)(D - b)y = F(x) = e^{cx} P_n(x); P_n(x) is a polynomial of degree n:

y_p = { e^{cx} Q_n(x) if c ≠ a and c ≠ b
x e^{cx} Q_n(x) if c = a or c = b, a ≠ b
x² e^{cx} Q_n(x) if c = a = b

Calculus of variations

$$J = \int_{x_1}^{x_2} f(y(x), y'(x); x)$$
$$\frac{\partial f}{\partial y} - \frac{d}{dx} \frac{\partial f}{\partial y'} = 0$$
$$\frac{\partial f}{\partial x} - \frac{d}{dx} \left(f - y' \frac{\partial f}{\partial y'} \right) = 0$$
$$f - y' \frac{\partial f}{\partial y'} = \text{const for } \frac{\partial f}{\partial x} = 0$$

FUNDAMENTAL (AND NOT SO FUNDAMENTAL) CONSTANTS	
$G =$	gravit. constant = $6.674 \times 10^{-11} \text{ N m}^2 \text{ kg}^{-2}$
$g =$	gravit. accel on Earth surface = 9.8 m/s^2
Note:	$GM_e = gR_e^2$
$R_S =$	mean radius of Sun = $696 \times 10^6 \text{ m}$
$R_E =$	mean radius of Earth = $6.371 \times 10^6 \text{ m}$
$R_M =$	mean radius of Moon = $1.737 \times 10^6 \text{ m}$
$R_{SE} =$	mean distance, Earth to Sun = $149.6 \times 10^9 \text{ m}$
$D =$	mean distance, Earth to Moon = $384.4 \times 10^6 \text{ m}$
$M_S =$	mass of Sun = $1.99 \times 10^{30} \text{ kg}$
$M_E =$	mass of Earth = $5.98 \times 10^{24} \text{ kg}$
$M_m =$	mass of Moon = $7.35 \times 10^{22} \text{ kg}$
PHYSICS	

Newton's laws

- 1st: Body remains at rest or in uniform motion unless acted upon by a force
- 2nd: $\mathbf{F}_{tot} = \frac{d\mathbf{p}}{dt} = m\mathbf{a}$
- 3rd: $\mathbf{F}_{A \rightarrow B} = -\mathbf{F}_{B \rightarrow A}$

Lagrangian dynamics

Hamilton's principle — Nature minimizes (makes stationary) the action.
Constrained — If a 3D system of N particles has $n < 3N$ minimum generalized coordinates, the system is *constrained*.
Natural — The coordinates q_i are *natural* if the relationships of r_α (every particle's position) to q_i doesn't change with time.
Ignorable — a coordinate q_i is *ignorable* if the corresponding generalized momentum p_i is constant.
Lyapunov Stability — If x_c is Lyapunov stable and all solutions that start out near x_c converge to x_c , then x_c is asymptotically stable

- Lagrangian:** $\mathcal{L} = T - U$
- Action:** $S = \int_{t_1}^{t_2} \mathcal{L}(q_1, q_2, \dots, q_N, \dot{q}_1, \dot{q}_2, \dots, \dot{q}_N, t) dt$
- Euler-Lagrange equations:** $\frac{\partial \mathcal{L}}{\partial q_1} = \frac{d}{dt} \frac{\partial \mathcal{L}}{\partial \dot{q}_1}, \dots$ etc.
- Generalized forces:** $F_i = \frac{\partial \mathcal{L}}{\partial q_i}$
- Generalized momenta** $p_i = \frac{\partial \mathcal{L}}{\partial \dot{q}_i}$

Orbits

- Definitions:** $M = m_1 + m_2$; $\mu = \frac{m_1 m_2}{m_1 + m_2}$; e.g., $U(\rho) = \frac{-Gm_1 m_2}{\rho}$
- Kinetic energy:** $T = \frac{1}{2} M \dot{r}^2 + \frac{1}{2} \mu \dot{r}^2$
- Lagrangian:** $\mathcal{L} = \frac{1}{2} \mu \dot{\rho}^2 + \frac{1}{2} \mu \rho^2 \dot{\phi}^2 - U(\rho)$
- Solution in ϕ :** $\dot{\phi} = \frac{L}{\mu \rho^2}$ (L const — angular momentum)
- Solution in ρ :** $\mu \ddot{\rho} = -\frac{d}{d\rho} U(\rho) + \frac{L^2}{\mu \rho^3} = -\frac{d}{d\rho} \left[U(\rho) + \frac{L^2}{2\mu \rho^2} \right]$
- Effective potential:** $U_{eff} = U(\rho) + \frac{L^2}{2\mu \rho^2}$
- Note cons. of energy:** $\frac{d}{dt} \left(\frac{1}{2} \mu \dot{\rho}^2 \right) = -\frac{d}{dt} U_{eff}(\rho)$; $E = \frac{1}{2} \mu \dot{\rho}^2 + U_{eff}(\rho)$
- Use:** $u = 1/r$ and $\frac{d}{dt} = \frac{d\phi}{dt} \frac{d}{d\phi} = \dot{\phi} \frac{d}{d\phi} = \frac{L}{\mu \rho^2} \frac{d}{d\phi} = \frac{L u^2}{\mu} \frac{d}{d\phi}$
- u-equation:** $u''(\phi) = -u(\phi) - \frac{\mu}{\ell^2 u(\phi)^2 F(u)}$

- Use:** $\gamma = Gm_1 m_2$ and $F(u) = -\gamma u^2$; then $U''(\phi) = -u(\phi) + \gamma \mu / \ell^2$; use $w(\phi) = u(\phi) - \gamma \mu / \ell^2$, so $W(\phi) = A \cos(\phi - \delta)$ ergo $u(\phi) = \frac{\gamma \mu}{\ell^2} + A \cos \phi$
- Radial eqn:** $r(\phi) = \frac{r_c}{1 + \epsilon \cos \phi}$
- Cartesian:**

$$\left(x + \frac{r_c \epsilon}{1 - \epsilon^2} \right)^2 + \left(\frac{y}{\frac{r_c}{\sqrt{1 - \epsilon^2}}} \right)^2$$

- Eccentricity:** $\epsilon = A \cdot r_c$ (A some constant)
- Circular orbit:** $r_c = \ell^2 / \gamma \mu$
- Min radius:** $r_{min} = \frac{r_c}{1 + \epsilon} = \frac{\ell^2}{\gamma \mu (1 + \epsilon)}$ (at $\phi = 0$; **perihelion**); $\ell = \mu r v_{tan}$
- Max radius:** $r_{max} = \frac{r_c}{1 - \epsilon}$ (at $\phi = \pi$; **aphelion**)
- Radial velocity:** $v_r = \sqrt{\frac{\mu}{r_c}} \cdot \epsilon \cdot \sin \phi$
- Tangential velocity:** $v_t = \sqrt{\frac{\mu}{r_c}} \cdot (1 + \epsilon \cdot \cos \phi)$
- Ellipse params:** $a = \frac{r_c}{1 - \epsilon^2}$; $b = \frac{r_c}{\sqrt{1 - \epsilon^2}}$; $d = a\epsilon$; $\epsilon = \sqrt{1 - (b/a)^2}$
- Orbital period:** $\tau = 2\pi \cdot a \cdot \sqrt{\frac{a}{\mu}}$
- Energy:** $E = \frac{\gamma^2 \mu}{2\ell^2} (\epsilon^2 - 1)$
- Kepler's 1st law:** Orbits are ellipses w/ sun at a focus (approx. true)
- Kepler's 2nd law:** Line from Sun to planet, equal areas in equal times
 $dA = \frac{1}{2} r^2 d\phi$; $\frac{dA}{dt} = \frac{1}{2} \frac{\ell}{\mu}$, ineq. of time
- Kepler's 3rd law:** $\tau = \frac{A}{dA/dt} = \frac{2\pi a b \mu}{\ell} \Rightarrow \tau^2 = 4\pi^2 \frac{a^3 r_c \mu^2}{\ell^2} = 4\pi^2 \frac{a^3 \mu}{\gamma} \approx \frac{4\pi^2}{GM_S} a^3$

Cartesian system

- inertia:** $I = \int_V r^2 dM$, $dM = \rho(x, y, z) dx dy dz$
- momentum:** $\mathbf{p} \equiv m\mathbf{v}$
- kinetic energy:** $T = \frac{1}{2} m v^2$

Rotating system

- moment of inertia:** $I = \int r^2 dm$ or, for a point mass, $I = r^2 m$, where r is the perp. distance to axis of rotation
- parallel axis theorem:** $I_z = I_{cm} + m d^2$; I_{cm} : inertia about center of mass, m : mass, d : distance between axes
- angular momentum:** $\mathbf{L} \equiv \mathbf{r} \times \mathbf{p} = I \boldsymbol{\Omega}$; \mathbf{r} : position vec, \mathbf{p} : linear momentum
- torque:** $\mathbf{N} \equiv \mathbf{r} \times \mathbf{F} = \dot{\mathbf{L}} = \mathbf{r} \times \dot{\mathbf{p}}$
- work:** $W = N\theta$, θ in rad
- angle:** θ
- angular velocity:** $\boldsymbol{\Omega} = \dot{\theta} = \frac{\mathbf{r} \times \mathbf{v}}{|\mathbf{r}|^2}$
- linear velocity:** $\mathbf{v} = \boldsymbol{\Omega} \times \mathbf{r}$
- angular acceleration:** $\boldsymbol{\alpha} = \dot{\boldsymbol{\Omega}} = \ddot{\theta} = \mathbf{a}_T / r$; \mathbf{a}_T is tangential acceleration
- newton's 2nd-law:** $\mathbf{N} = I \boldsymbol{\alpha}$
- time deriv, unit vec in rotating frame:** $\left(\frac{d\mathbf{e}}{dt} \right)_{S_0} = \left(\frac{d\mathbf{e}}{dt} \right)_S + \boldsymbol{\Omega} \times \mathbf{e}$ (S_0 : inert, S : rot)
- Newton's 2nd in rotating frame:** $m \ddot{\mathbf{r}} = \mathbf{F} + 2m \dot{\mathbf{r}} \times \boldsymbol{\Omega} + m (\boldsymbol{\Omega} \times \mathbf{r}) \times \boldsymbol{\Omega}$

Conservative Force

- Conditions, given that \mathbf{F} has continuous first partial derivatives in a simply connected region. . .
- No curl anywhere:** $\nabla \times \mathbf{F} = 0$
- Equal work regardless of path :**
 $W_C = \int_C \mathbf{F} \cdot d\mathbf{s} = \text{constant} \forall C$
 $W_C = \oint_C \mathbf{F} \cdot d\mathbf{s} = 0 \forall C$
- $\mathbf{F} \cdot d\mathbf{r}$ is exact differential**
 $\mathbf{F} = \nabla W$, W single-valued
- Allows definition of potential:** $\mathbf{F} = -\nabla U$

Specific Forces

- gravity :**
 - point mass or sph.-symm mass: $\mathbf{F} = -G \frac{Mm}{r^2} \mathbf{e}_r \approx -mg$ on earth
 - generally: $\mathbf{F} = -Gm \int_V \frac{\rho(\mathbf{r}') \mathbf{e}_r}{r^2} dv'$
 - grav field vector: $\mathbf{g} \equiv -\nabla \Phi = \mathbf{F}/m$
 - grav potential, point mass: $\Phi = -G \frac{M}{r}$
 - grav potential, mass distr: $\Phi = -G \int_V \frac{\rho(\mathbf{r}')}{r} dv'$
 - potential energy: $U = m\Phi$
 - Gauss' law for grav, int: $\oint_S \mathbf{g} \cdot d\mathbf{A} = -4\pi GM$
 - Gauss' law for grav, dif: $\nabla \cdot \mathbf{g} = -4\pi G\rho$
 - Poisson's equation: $\nabla^2 \phi = 4\pi G\rho$, for rad-sym system, this is $\frac{1}{r^2} \frac{\partial}{\partial r} \left(r^2 \frac{\partial \phi}{\partial r} \right) = 4\pi G\rho(r)$ and $\mathbf{g}(r) = -\mathbf{e}_r \frac{\partial \phi}{\partial r}$
- tidal** (due to Moon's gravity)
 - \mathbf{e}_R points from Moon's center to test mass on Earth
 - \mathbf{e}_D is from center of Moon to center of Earth
 - (x, y) ECEF coord of test mass
 - $\mathbf{F}_T = -GmM_m \left(\frac{\mathbf{e}_R}{R^3} - \frac{\mathbf{e}_D}{D^3} \right)$
 - $F_{Tx} \approx 2GmM_m x / D^3$
 - $F_{Ty} \approx -GmM_m y / D^3$
- spring** (simple, linear)
 - $F = -kx$ (x : displ from eq lib, k : spring const)
 - $U = \frac{1}{2} k x^2$
- inertial force, linear accel:** $\mathbf{F}_{inert} = -m\mathbf{A}$ (\mathbf{A} : frame's accel w.r.t. inertial frame)
- centrifugal** (inertial force)
 - $\mathbf{F}_{centr} = m (\boldsymbol{\Omega} \times \mathbf{r}) \times \boldsymbol{\Omega}$ (generally)
 - $\mathbf{F}_{centr} = \frac{mv^2}{r} \mathbf{e}_r = m r \Omega^2 \mathbf{e}_r$ (for circular motion)
 - $U_{centr}(r) = \frac{\ell^2}{2mr^2}$ (ℓ : angular momentum)
 - Free-fall accel (e.g., on Earth): $\mathbf{g} = \mathbf{g}_0 + (\boldsymbol{\Omega} \times \mathbf{R}) \times \boldsymbol{\Omega}$
- coriolis** (inertial force)
 - $\mathbf{F}_{cor} = 2m \dot{\mathbf{r}} \times \boldsymbol{\Omega}$
- buoyant** —
 - $F = \rho_{fluid} V g$, dir. opposite to grav.-induced pressure grad. in fluid;
 - ρ_{fluid} : density, V : submerged volume, g : grav.
- lorentz** —
 - $\mathbf{F} = q \mathbf{v} \times \mathbf{B}$, q is charge of particle, \mathbf{v} its velocity, \mathbf{B} is mag. field strength
- electrostatic** —
 - $\mathbf{F} = q\mathbf{E}$, q is charge, \mathbf{E} electric field

Energy

- potential energy:** $\int_1^2 \mathbf{F} \cdot d\mathbf{r} \equiv U_1 - U_2$
(work, done by force \mathbf{F} , req'd to move particle from point 1 to point 2 with no change in kinetic energy); potential energy is the capacity to do work.
- force due to the potential U :** $\mathbf{F} = -\nabla U$
- kinetic energy:** $T \equiv \frac{1}{2} m v^2$
- total energy:** $E \equiv T + U$
- 1D solution given E and $U(x)$,** for conservative force only:

$$t - t_0 = \int_{x_0}^x \frac{\pm dx}{\sqrt{\frac{2}{m} [E - U(x)]}}$$

Conservation theorems

- linear momentum:** $\frac{d}{dt} (p_1 + p_2) = 0$ (or $p_1 + p_2$ is const) if no external forces act upon system
- angular momentum:** $\dot{\mathbf{L}} = \mathbf{r} \times \dot{\mathbf{p}} = 0$ (or \mathbf{L} is const) if no external torque acts upon system
- energy:** $\mathbf{F} + \nabla U = 0$; $\frac{dE}{dt} = 0$ if the force field represented by \mathbf{F} is conservative

Harmonic oscillation

Simple harmonic oscillator

$$\begin{aligned} m\ddot{x} &= -kx \\ \omega_0^2 &\equiv k/m \\ \ddot{x} + \omega_0^2 x &= 0 \\ x(t) &= A \sin(\omega_0 t - \delta) \\ E &= T + U = \frac{1}{2} k A^2 \end{aligned}$$

Damped oscillator

equation of motion: $m\ddot{x} + b\dot{x} + kx = 0$; b is resisting force coeff, k is restoring force coeff
convenient substitutions: $\beta \equiv \frac{b}{2m}$ (damping), and $\omega_0^2 \equiv k/m$ (natural ang. freq, undamped sys)
new eqn of motion: $\ddot{x} + 2\beta\dot{x} + \omega_0^2 x = 0$
general sol'n:

$$x(t) = e^{-\beta t} \left[A_1 \exp\left(\sqrt{\beta^2 - \omega_0^2} t\right) + A_2 \exp\left(-\sqrt{\beta^2 - \omega_0^2} t\right) \right]$$

underdamping: $\omega_0^2 > \beta^2$
critical damping: $\omega_0^2 = \beta^2$
overdamping: $\omega_0^2 < \beta^2$

Sinusoidally-driven damped oscillator

eqn of motion: $m\ddot{x} + b\dot{x} + kx = F_0 \cos \omega t$; b is resisting force coeff, k is restoring force coeff
convenient substitutions: $A = F_0/m$ (driving ampl), $\beta \equiv \frac{b}{2m}$ (damping), and $\omega_0^2 \equiv k/m$ (natural ang. freq, undamped sys)
new eqn of motion: $\ddot{x} + 2\beta\dot{x} + \omega_0^2 x = A \cos \omega t$

complementary solution:

$$x_c(t) = e^{-\beta t} \left[A_1 \exp\left(\sqrt{\beta^2 - \omega_0^2} t\right) + A_2 \exp\left(-\sqrt{\beta^2 - \omega_0^2} t\right) \right]$$

particular solution:

$$\begin{aligned} x_p(t) &= \frac{A}{\sqrt{(\omega_0^2 - \omega^2)^2 + 4\omega^2 \beta^2}} \cos(\omega t - \delta) \\ \delta &= \arctan\left(\frac{2\omega\beta}{\omega_0^2 - \omega^2}\right) \end{aligned}$$

amplitude resonance frequency: $\omega_R = \sqrt{\omega_0^2 - 2\beta^2}$ ($\omega_r < \omega_1 < \omega_0$)

kinetic energy resonance frequency: $\omega_E = \omega_0$

quality factor: $Q \equiv \frac{\omega_R}{2\beta} \approx \frac{\omega_0}{\Delta\omega}$ (the latter is for lightly damped systems; $\Delta\omega$ is the distance between half-energy points — $D_{res}/\sqrt{2}$ — on the amplitude resonance curve)

Underdamped oscillator

pseudo-frequency of oscillation: $\omega_1^2 \equiv \omega_0^2 - \beta^2$

solution (form 1): $x(t) = e^{-\beta t} \left[A_1 e^{i\omega_1 t} + A_2 e^{-i\omega_1 t} \right]$

solution (form 2): $x(t) = A e^{-\beta t} \cos(\omega_1 t - \delta)$

phase plot: Use the var. subst. $u = \omega_1 x$, $w = \beta x + \dot{x}$ and plot w on the y -axis vs. u on the x -axis

response to δ force: $x(t) = \frac{b}{\omega_1} e^{-\beta(t-t_0)} \sin \omega_1(t - t_0)$

green's fn: $G(t, t') \equiv \frac{1}{m\omega_1} e^{-\beta(t-t')} \sin \omega_1(t - t')$, $t \geq t'$; 0 otherwise

Critically damped oscillator

qualitative behavior: System approaches equilibrium (natural solution dies out) faster than the others.

solution: $x(t) = (A + Bt)e^{-\beta t}$

Overdamped oscillator

pseudo-frequency of (non-)oscillation: $\omega_2^2 \equiv \beta^2 - \omega_0^2$

solution: $x(t) = e^{-\beta t} [A_1 e^{\omega_2 t} + A_2 e^{-\omega_2 t}]$

phase plot: Asymptotic behavior tends towards $\dot{x} = -(\beta - \omega_2)x$ unless $A_1 = 0$, then it goes to $\dot{x} = -(\beta + \omega_2)x$

Series RLC circuit

voltage across inductor: $V_L = L \frac{dI}{dt} = L \ddot{q}$

voltage across resistor: $V_R = IR = R \frac{dq}{dt} = R \dot{q}$

voltage across capacitor: $V_C = \frac{q}{C}$

diff eq of RLC circuit with driving power source: $L\ddot{q} + R\dot{q} + q/C = V(t)$

Electrical–mechanical equivalents

	Mechanical		Electrical
x	Displacement	q	Charge
\dot{x}	Velocity	$\dot{q} = I$	Current
m	Mass	L	Inductance
b	Damping resistance	R	Resistance
$1/k$	Mechanical compliance	C	Capacitance
F	Ampl of impressed force	\mathcal{E}	Ampl of impressed emf