

SDSS-IV MANGA: THE UNIVERSAL PROFILE OF ENHANCED STAR FORMATION IN CLOSE GALAXY PAIRS

JOSHUA L. STEFFEN¹, HAI FU¹, AND MANGA TEAM

Draft version August 28, 2020

ABSTRACT

We study the spatial profile of the star formation rate (SFR) in a sample of 220 interacting galaxies within the SDSS-IV MaNGA survey. We find that the specific star formation rate (sSFR) is centrally (within one effective radius) enhanced in interacting galaxies by $\sim 0.25 \pm 0.1$ but that there is no enhancement or suppression in their disks. We split the radial profile into five separate stellar mass bins from $\log(M/M_\odot) = 9.0$ – 11.5 and find that the sSFR enhancement has no dependence on the stellar mass of the paired galaxies. Within one effective radius, the sSFR enhancement in each of the mass bins falls within ~ 0.1 dex of the median profile for the whole pair sample. We then use a second method to compare paired galaxies and control galaxies where each paired galaxy is matched to a set of control galaxies based on stellar mass and redshift. The sSFR enhancements found are still mass independent, reinforcing the idea that low mass galaxies experience the same level of enhancement to their star formation rate as high mass galaxies. We describe the profiles of the $\Delta\log(\text{sSFR})$ enhancement as a universal profile because the general shape of the profile is unchanged by parameters like stellar mass, mass ratio, or projected separation. The shape of the profile is a central peak in $\Delta\log(\text{sSFR})$ which linearly falls to zero around $1.0 R/R_{\text{eff}}$. Parameters like the mass ratio or projected separation may scale the height of the profile’s central peak, but the general shape of the profile is preserved.

Keywords: galaxies: star formation — galaxies: nuclei — galaxies: interactions — galaxies: mass evolution

1. INTRODUCTION

In the Λ CDM model, galaxy evolution is a hierarchical process. In this model, massive galaxies are the product of several past merger events of smaller galaxies. In fact, cosmological hydrodynamical simulations have shown that repeated merger events may be responsible for as much as $\sim 60\%$ of stellar mass in massive galaxies like M87 (e.g. [Rodríguez-Gomez et al. \(2016\)](#); [Pillepich et al. \(2018\)](#)).

The internal dynamics of these interacting galaxies were first modeled in the seminal work, [Toomre & Toomre \(1972\)](#). Since then, hydrodynamical simulations have expanded upon the N-body simulations of [Toomre & Toomre \(1972\)](#) by modeling gas-dynamics within the galaxies. These simulations show how barred structures develop within the disks of the interacting galaxies due to the tidal torques between them ([Barnes & Hernquist 1991](#)). As the bars form, the gases within the galaxy’s disk lose angular momentum and get funneled into the centers of the galaxies.

When the gas-inflows impact upon the gases in the nucleus of galaxy a burst of new star formation is triggered ([Barnes & Hernquist 1996](#); [Mihos & Hernquist 1996](#)). These gas inflows will also bring metal-poor gases from the disk into the center of the galaxy which can dilute the central metallicity ([Rupke et al. 2010](#); [Perez et al. 2011](#); [Scudder et al. 2012](#)). The gas-inflows may also be able to make it into very center of the galaxy and trigger supermassive black hole (SMBH) accretion ([Capelo et al. 2017](#)).

Centrally enhanced star formation has also been observed in galaxy pairs with large spectroscopic surveys including; CfA2 (Center of Astrophysics Redshift Survey) ([Barton et al. 2000](#); [Woods et al. 2006](#)), 2dF (Two Degree Field) ([Lambas et al. 2003](#)), AEGIS (All-Wavelength Extended Groth Strip International Survey) ([Lin et al. 2007](#)), SDSS (Sloan Digital Sky Survey) ([Ellison et al. 2008](#)), COSMOS (Cosmic Evolu-

tion Survey) ([Kartaltepe et al. 2007](#); [Xu et al. 2012](#)), PRIMUS (Prism Multi-Object Survey) ([Wong et al. 2011](#)), and GAMA (Galaxy and Mass Assembly) ([Robotham et al. 2014](#)).

This star formation enhancement increases with closer projected separations between the two galaxies ([Li et al. 2008](#); [Ellison et al. 2008](#); [Scudder et al. 2012](#)) but has still shown to be detectable out to projected separations of 150 kpc ([Patton et al. 2013](#)). The level of the star formation enhancement also increases with smaller stellar mass ratios, $|\Delta\log(M)|$, between the two galaxies ([Ellison et al. 2008](#)). Similar mass galaxies show the strongest sSFR effects while large central galaxies interacting with small satellite galaxies show weaker sSFR effects. Because of this, most surveys will split their sample into major mergers, $|\Delta\log(M)| \leq 0.5$, and minor mergers, $0.5 \leq |\Delta\log(M)| \leq 1.0$, where major mergers are expected to show the strongest interaction induced effects.

With the recent large integral field spectroscopic (IFS) surveys, interacting galaxies can be studied with unprecedented spatial detail. These surveys allow astronomers to study the centers of merging galaxies more rigorously since apertures can be set to the physical scale of the galaxies instead of being bound by a fixed sky aperture. These IFS surveys will also allow us to see the extent of the centrally induced star formation and to see how the star formation in the disks of the galaxies are affected.

The IFS surveys, CALIFA² and MaNGA³, have already been used to study the SFR of interacting galaxies in merging galaxies. The SFR was found to be centrally enhanced by a factor of 2 – $3\times$ ([Barrera-Ballesteros et al. 2015](#); [Pan et al. 2019](#)), which is in agreement with the previous single-fiber spectroscopic surveys. These surveys also studied the SFR in the disk of the paired galaxies; however, there is a standing discrepancy between the different works. [Barrera-Ballesteros et al. \(2015\)](#) found a suppression to the SFR of $\sim 0.74\times$ while [Pan et al. \(2019\)](#) found an enhancement of 1.22 – $1.58\times$.

¹ Department of Physics & Astronomy, The University of Iowa, 203 Van Allen Hall, Iowa City, IA 52242

² Calar Alto Legacy Integral Field Area

³ Mapping Nearby Galaxies at Apache Point Observatory

The MaNGA survey has also been used to study the radial profiles of SFR in post-merger galaxies, which are the galaxies which have recently coalesced. Hydrodynamical simulations predict a second burst of star formation as the two galaxies coalesce. This effect has been observed in a sample of post-merger galaxies which showed a central enhancement to the SFR surface density of $2.5\times$ and an enhancement in their disks of $1.26\times$ (Thorp et al. 2019).

In our previous work, Fu et al. (2018), we built a sample of paired galaxies in the MaNGA survey where both components of the pair were contained within field of view of a single integral field unit (IFU). We found that approximately 5.7% of the MaNGA galaxies have a companion galaxy contained within the field-of-view of a single IFU. In this work, we will use a version of this sample updated for the recent MaNGA Data Release 16 (DR16) (Ahumada et al. 2020) and we will further supplement this with a sample of companion galaxies identified outside the field-of-view of the MaNGA IFUs.

In this work we will study how the radial profiles of star formation in our pair sample is effected by tidal interactions. We will expand upon the previous studies of star formation in merging galaxies by studying how the spatial profile of the SFR behaves as a function of the stellar mass of the paired galaxy, the mass ratio between the galaxy pairs, and the projected separation between the galaxies.

Changes in the star formation rate will also affect the future mass growth in the interacting galaxies. We will model the fractional mass growth rate of the galaxies in our sample to see how the merger event influences the mass evolution of the paired galaxies. Since merging galaxies are already known to have centrally enhanced star formation rates, we would expect this to produce higher concentrations of stellar mass in their centers.

This paper is organized as follows; in Section 2 we will discuss the properties of the MaNGA survey along with the construction of our pair and control samples, in Section 3 we will discuss how we measure star formation rates and how we build radial profiles of star formation, in Section 4 we discuss the mass independence of the constructed radial profiles, in Section 5 we explore the radial profiles as a function of the mass ratio and projected separation and we compare our work against previous works, and in Section 6 we summarize the findings of the work.

2. DATA AND SAMPLE

MaNGA is an IFS survey based at Apache Point Observatory (APO) which used the SDSS (Sloan Digital Sky Survey) 2.5-meter telescope along with two dual-channel BOSS spectrographs (Drory et al. 2015). MaNGA captures spectra through 17 integral field units (IFUs) with variable numbers of fibers; 19, 37, 61, 91, and 127 fibers covering $12.5''$, $17.5''$, $22.5''$, $27.5''$, and $32.5''$ on the sky respectively Law et al. (2015). MaNGA is an optical survey with a spectral coverage of $3600\text{--}10,300\text{ \AA}$ with a resolution of $R \sim 2000$ and a median spectra coverage of $2.5''$ FWHM (Bundy et al. 2015).

The MaNGA survey targets galaxies from a subset of 41,154 galaxies from the NASA-Sloan Atlas (NSA v1_0_1; <http://www.ns atlas.org>) with a redshift range of $0.01 < z < 0.15$ and a luminosity range of $-17.7 < \mathcal{M}_i < -24.0$, where \mathcal{M}_i is the rest frame i-band magnitude within the survey’s elliptical Petrosian apertures. MaNGA plans to cover 10,000 galaxies with a flat stellar mass distribution at

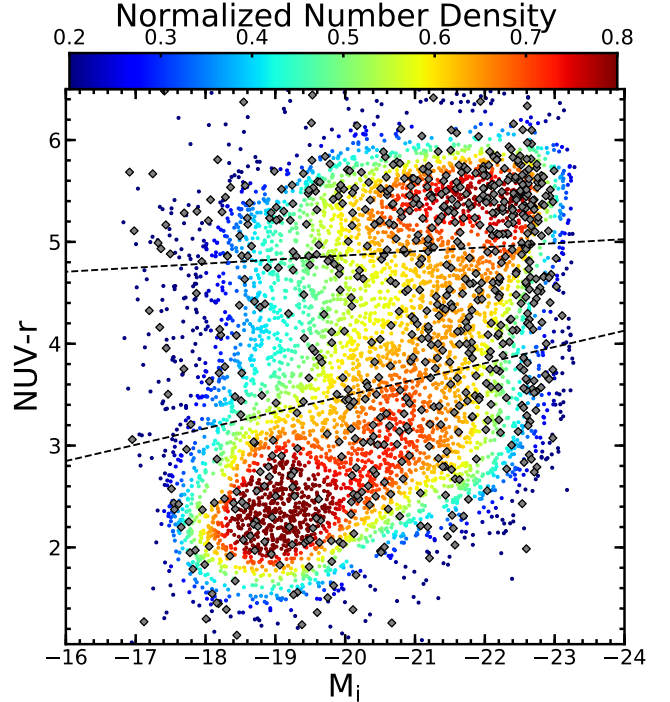


Figure 1. Color-magnitude diagram for MaNGA galaxies (colored circles). The color of the symbol reflects the local density around each data point in this color-magnitude plane, as indicated by the color bar on the top. Inside and outside-IFU pairs are marked with grey diamonds. From top to bottom, the dashed lines divide the sample into red sequence, green valley, and blue cloud. The star-forming galaxy sample used in this paper are the galaxies below the lower dividing line.

two spatial coverages, $1.5 R/R_{\text{eff}}$ and $2.5 R/R_{\text{eff}}$ (where R_{eff} is the radius which contains 50% of the galaxy’s total light). In this work we use the data from the internal data release, MPL-8, which covers 6142 unique galaxies.

We classify galaxies in this survey as star forming galaxies using the color-magnitude diagram (CMD). We show the CMD for the MaNGA survey and our pair sample in Figure 1 along with demarcation lines which separate the blue cloud, red sequence, and green valley. We established the demarcation lines by collapsing the CMD to a color histogram for each of the three regions. We then varied the slopes between the regions until we found the slopes which best fit the data. These demarcation lines are;

$$NUV-r = 3.1682 - 0.16(\mathcal{M}_i + 18) \quad (1)$$

$$NUV-r = 4.7866 - 0.04(\mathcal{M}_i + 18) \quad (2)$$

Where $NUV-r$ is the color and \mathcal{M}_i is the i-band magnitude from the NSA catalog. We have checked the selected star forming galaxies on the BPT diagram (Baldwin et al. 1981) and confirmed that there is not a significant number of AGN in our sample.

On top of the color selection cuts, we require that all galaxies are in the Primary or Secondary MaNGA subsamples (Wake et al. 2017), and that the stellar mass range of the galaxy sample is between $\log(M/M_{\odot}) = 9.0\text{--}11.5$.

We build two different pair samples in this work; the inside-IFU sample which contains paired galaxies where both galaxies are covered by a single MaNGA IFU and the outside-IFU sample where a MaNGA target galaxy is coupled with another

galaxies found outside of its MaNGA IFU.

2.1. Inside-IFU Sample

In our previous work, [Fu et al. \(2018\)](#), using the 2168 unique galaxies from DR14 ([Abolfathi et al. 2018](#)), we showed that $\sim 5.7\%$ of the MaNGA observations have at least one companion galaxy within the field of view of the IFU. In this work we use a refined pair identification process which is also expanded for the most recent internal data release, MPL-8.

We identify potential paired galaxies by overlaying SDSS photometric objects over the MaNGA fields of view. We manually inspect each MaNGA field, removing photometric objects which are over-deblended galaxy fragments and adding any objects missed by the SDSS photometric catalog. At this stage the object catalog includes any foreground stars and foreground/background galaxies along with the potential paired galaxies.

To select paired galaxies out of our object catalog, we inspect the spectra of each object. We use our own IDL based spectral fitting code, SPFIT, to model the spectra from the MaNGA datacubes. SPFIT simultaneously fits emission lines and the stellar continuum with the Levenberg-Marquardt non-linear least-squares minimization algorithm ([Fu et al. 2018](#)). Emission lines are parameterized as a Gauss-Hermite series and the stellar continuum is the sum of simple stellar populations (SSPs) convolved with the line-of-sight velocity distribution (LOSVD).

The spectra of the identified objects is extracted through a $1''$ circular aperture and fitted assuming the MaNGA target’s redshift and then manually sorted into the following categories; “good” galaxy spectra, broad-line AGN (active galactic nuclei), foreground star, foreground/background galaxies, or poor S/N objects. The “good” galaxy spectra are the objects whose spectra is well modeled by SPFIT at the target galaxy’s redshift, whether it is the target galaxy itself or a nearby companion galaxy. This means that the companion galaxy can be within $\pm 3000 \text{ km s}^{-1}$ of the MaNGA target. We found 6573 “good” objects, 57 broad-line AGN, 836 foreground stars, 319 foreground/background galaxies, and 1546 objects with poor S/N.

From the 6573 galaxies with good spectra, 404 of the MaNGA IFUs have multiple objects within the IFU. We further restrict the sample by setting a relative velocity cut of $\Delta v < 500 \text{ km s}^{-1}$ to remove projected companions. Again, the galaxy sample is also restricted to galaxies with a stellar mass range of $\log(M/M_{\odot}) = 9.0\text{--}11.5$ and galaxies in the Primary or Secondary MaNGA subsamples. We also require that the galaxies are classified as star forming by the CMD diagram. These requirements reduce our sample down to 87 star forming galaxies (SFGs).

Several of the target galaxies in our sample have multiple companions. From the whole set of 404 IFUs, there are 327 pairs, 67 triplets, 7 quadruplets, and 1 quintuplet.

2.2. Outside-IFU Sample

We supplement the inside-IFU sample with a set of pairs identified outside of the field of view of the MaNGA IFU. We select these outside-IFU pairs from the NSA catalog. We search for these external pairs by selecting objects with a projected separation from the MaNGA targets of $r_p < 50 \text{ kpc}$ using the MaNGA target’s redshift. We further use a relative velocity cut of $\Delta v < 500 \text{ km s}^{-1}$ to remove projected galaxies from the selection.

From the NSA catalog’s 641,409 galaxies, we find 492 galaxies which are paired to MaNGA targets. MaNGA targets which have both an inside-IFU and an outside-IFU pair are left to the inside-IFU sample. After restricting the sample to SFGs, we have another 133 MaNGA targets with paired galaxies outside of the IFU. This gives us a total of 220 MaNGA targets in our pair sample.

2.3. Control Sample

The SFR in galaxies has been shown to be dependent on both stellar mass and redshift of the galaxies ([Noeske et al. 2007](#)). This effect needs to be controlled for when comparing the paired and control galaxies to isolate the changes in the sSFR due to merger induced effects.

We build a sample of isolated control galaxies from the remaining MaNGA galaxies by selecting the MaNGA target galaxies which have no spectroscopic companions within $r_p < 50 \text{ kpc}$ and $\Delta v < 500 \text{ km s}^{-1}$ either inside or outside of the IFU. This gives us a set of 1891 star forming control galaxies to work with.

We then set two separate methods for comparing the paired galaxies with the control galaxies. In the first method, we simply split the galaxies into separate stellar mass bins. We then compare paired galaxies and control galaxies within each stellar mass bin.

In the second method, we control for both stellar mass and redshift. To do this, we set a 0.1 dex stellar mass limit between a paired galaxy and the potential control galaxies. Then we restrict the redshift between the control and paired galaxies, by selecting control galaxies which are drawn from the same MaNGA subsample (e.g. Primary or Secondary) as the paired galaxy. The MaNGA sample has a tight “banana” shaped distribution in stellar mass and redshift. By setting a 0.1 dex stellar mass cut and requiring that the galaxies be in the same MaNGA subsample, we effectively set a $z \leq \sim 0.025$ redshift cut.

Setting the above requirements, most paired galaxies will find between 20–100 control galaxies. Since we want each paired galaxy to be treated in a similar manner, we will downselect the total number of acquired control galaxies down to 20 controls using a random number generator. In the cases where a paired galaxy does not initially acquire 20 control galaxies we iteratively expand the stellar mass limit by 0.1 dex until at least 20 control galaxies are found.

A small set of paired required extra iterations to find 20 controls; 8 paired galaxies need an extra iteration, 1 paired galaxy need two extra iterations, and 1 paired galaxy needed three iterations.

While the first method is very basic in comparison to the second method, we will see in Section 4 that the first method is sufficient to see the central enhancement to the star formation rate and the mass independence of the star formation enhancement.

3. RADIAL PROFILES OF STAR FORMATION

3.1. Emission Line Extraction

We use SPFIT to get measurements of emission line fluxes, equivalent widths, velocities, and velocity dispersions of 19 emission lines and measurements of stellar masses, ages, [Fe/H], and the kinematics of the stellar populations.

We correct the emission lines for reddening using the extinction curve from [Cardelli et al. \(1989\)](#) with updated coefficients from [O’Donnell \(1994\)](#). The extinction in parameter-

ized as $R_V \equiv A_V/E(B-V) = 3.1$, where we estimate the value of the V-band extinction, A_V , by comparing the $H\alpha/H\beta$ ratio to the expected value of 2.85 for case-B recombination.

$$A_V = 6.78 \log [(H\alpha/H\beta)/2.85] \quad (3)$$

The V-band extinction is then used to calculate the wavelength dependent extinction A_λ :

$$A_\lambda(x) = A_V (a(x) + b(x)/R_V) \quad (4)$$

where $x = 1/\lambda$ and $a(x)$ and $b(x)$ are the coefficients for the extinction curve from O'Donnell (1994). The extinction corrected flux is then;

$$F'(\lambda) = F(\lambda) * 10^{0.4A_\lambda(x)} \quad (5)$$

This extinction correction is applied to all of the extracted emission lines.

3.2. Star Formation Rate

We measure the SFR from the $H\alpha$ luminosity, $L_{H\alpha}$, calculated from the $H\alpha$ flux. We use the SFR formula, Equation 6, from Murphy et al. (2011) which uses a Kroupa IMF, Solar metallicity, a constant SFR at an age of 100 Myr, and Case-B recombination.

$$\frac{\text{SFR}}{M_\odot \text{ yr}^{-1}} = \frac{L_{H\alpha}}{1.86 \times 10^{41} \text{ erg s}^{-1}} \quad (6)$$

Since the stellar mass of a galaxy is not uniformly distributed within the galaxy, we normalize the SFR by the local stellar mass, giving us the specific star formation rate (sSFR). The local stellar masses used here is the stellar mass calculated by SPFIT's fit to the stellar continuum within each spaxel (spatial pixel).

$$\text{sSFR} = \frac{\text{SFR}}{M^*} \quad (7)$$

This will allow us to compare the SFR in high mass regions, like the centers of the galaxies, to the SFR lower mass regions, like the disk of the galaxies. We have also found a strong correlation between sSFR and $\text{EW}(H\alpha)$ and have found the same results using $\text{EW}(H\alpha)$. This is useful to check as $\text{EW}(H\alpha)$ is an observable and independent of SPFIT's measurement of the stellar mass.

It has been shown that hot low mass evolved stars (HOLMES) in retired galaxies can mimic H II regions and AGN on the BPT diagram (Stasińska et al. 2008). To mask retired spaxels, we implement a $H\alpha$ equivalent width (EW) cut of $\text{EW}(H\alpha) > 3\text{\AA}$ (Cid Fernandes et al. 2011).

3.3. Geometry of Galaxies

In order to spatially characterize the star formation in the paired galaxies we build radial profiles of sSFR. First, the geometry of the galaxies need to be defined. Specifically we will need position angle, inclination angle, and the effective radius of each of the MaNGA targets. We use the r -band elliptical Petrosian apertures from the NSA catalog and the r -band Sérsic apertures from Simard et al. (2011) to define the geometries of the galaxies.

The NSA catalog has complete coverage over the MaNGA sample (since MaNGA selects its targets from this catalog); however, it tends to fail to properly fit paired galaxies with

close on-sky separations. We found that the apertures from Simard et al. (2011) work better for these close galaxies, but the catalog does not completely cover the MaNGA sample. We use NSA catalog for the outside-IFU pair sample because they are well separated on the sky and the Simard et al. (2011) catalog for the inside-IFU sample. In the cases where an Inside-IFU paired galaxy is not covered by Simard et al. (2011), we use the ellipse from the NSA catalog.

We fit the geometry with aperture from both catalog when available. When using the first method of pair-control comparison, where pairs and controls are grouped into stellar mass bins, we fit the control galaxies with the NSA apertures. When making the comparison with the second method, paired galaxies fitted with the NSA aperture are compared against controls fitted with the NSA aperture and paired galaxies fitted with the Simard et al. (2011) apertures are compared against controls fitted with the Simard et al. (2011) apertures.

We compared the magnitudes extracted from the two apertures for the control galaxies and found that they are in close agreement. This means that the geometries between the two surveys must be in close agreement and the profiles of a galaxy fitted with a NSA aperture can be confidently compared against a galaxy fitted with an aperture from Simard et al. (2011).

We calculate the inclination angle, i , of the galaxies using the major-to-minor axis ratios from the elliptical apertures;

$$\cos^2(i) = \frac{(b/a)^2 - q^2}{1 - q^2} \quad (8)$$

Where q is the intrinsic oblateness. We use the empirically determined oblateness of $q = 0.13$, from Giovanelli et al. (1994).

The inclination angle, along with the galaxy's position angle, is used to deproject the geometries of the galaxies. We use the 50% half light radius, the effective radius (R_{eff}) to scale the sizes of the galaxies. Doing this will allow use the compare galaxies of different sizes against each other.

3.4. Radial Profiles

Once the geometry of the galaxies are set, we can build the radial profiles. The spaxels are binned into radius increments of $0.1 R/R_{\text{eff}}$ from $0.0-2.5 R/R_{\text{eff}}$. Within each radius bin we take the median of the specific star formation rate. The profiles' errors are the standard error of the mean of the data within each radius bin. We create one of these individual averaged radial profiles for every MaNGA galaxy.

We use our catalog of photometric objects and the Simard et al. (2011) catalog to mask foreground/background objects within the MaNGA IFUs which may contaminate the data. Objects in our photometric catalog are masked out with a $2''$ circular aperture and the galaxies from Simard et al. (2011) are masked out with a $2.0 R/R_{\text{eff}}$ elliptical aperture, using the r -band Sérsic apertures from the catalog.

4. RESULTS

4.1. Star Formation Enhancement

With individual $\log(\text{sSFR})$ profiles built for each MaNGA galaxy, we now use two different methods to compare the profiles of paired galaxies against the profiles of control galaxies. The first method, in Section 4.1.1, paired galaxies are compared to control galaxies within evenly spaced stellar mass-bins. In the second method, in Section 4.1.2, paired galaxies

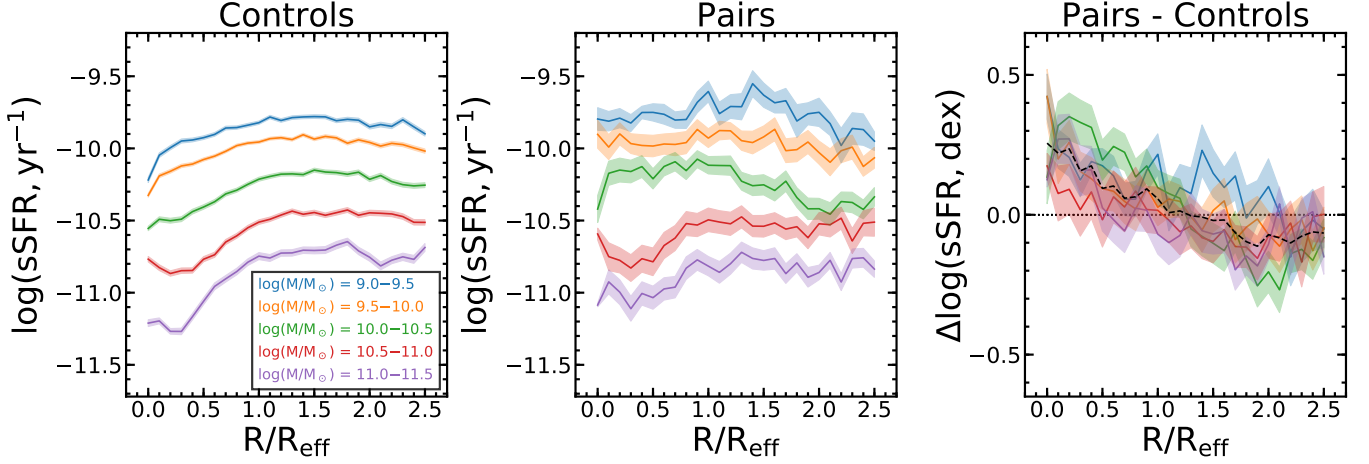


Figure 2. The $\log(\text{sSFR})$ as a function of galactocentric radius for control galaxies (**Left**) and galaxy pairs (**Middle**). The difference between the profiles of the paired galaxies and the control galaxies are shown in the **Right** panel. The dashed black profile represents the mean of the difference profiles). The colors of the profiles represent the mass range of the selected galaxies and the highlighted region around the profiles represent the standard error of the mean of the data at the given radius interval.

are compared to a subset 20 control galaxies which have a similar stellar mass and redshift.

4.1.1. Mass-Binned Difference Profiles

In the first method, paired and control galaxies are grouped into five evenly spaced stellar mass bins between $\log(M/M_\odot) = 9.0$ –11.5. Within each stellar mass bin, we create a median profile of $\log(\text{sSFR})$ as a function of galactocentric radius for both paired and control galaxies. The control profiles are constructed using all available control galaxies within each mass bin. The error associated with the profile is the standard error of mean of the data at each radius bin. We show these “stacked” profiles for control galaxies in the left hand panel of Figure 2 and the stacked profiles for paired galaxies in the middle panel of Figure 2.

Once the stacked profiles are made, we take the difference between the stacked profiles of the paired galaxies and the control galaxies, pair - control, in log space (this means that the profiles really represents a ratio between the pairs and controls in linear space). This gives us difference profile, $\Delta\log(\text{sSFR})$, which shows us where the sSFR is enhanced or suppressed (shown in the right hand panel of Figure 2).

This method has the benefit of a high S/N control profile since each mass bin will have hundreds of control galaxies while the median control profiles of the second method will only have 20 controls. On the other hand, this method only controls for a single galaxy parameter, stellar mass. Further the profiles can only be studied as a function of stellar mass so we are unable to study the profiles as a function of projected separation or mass ratio.

We show the mass-binned profiles in Figure 2. Between the median profiles of the control galaxies and paired galaxies, we can see that the paired profiles are flatted in their centers in comparison to the control profiles. The paired profiles are centrally enhanced by $\sim 0.25 \pm 0.1$ dex. This enhancement linearly falls to zero around $1.0 R_{\text{eff}}$. In the disks the profiles feature a weak suppression to the sSFR of 0.05–0.1 dex.

The sSFR enhancement show no dependence on the stellar mass, the profiles of the different mass bins fall within ~ 0.2 dex of median profile. This means that low mass galaxies are enhanced at the same level as high mass galaxies. It is also notable that while this method of pair-control comparison is

not rigorous as the next method, the central star formation enhancement is clearly present.

4.1.2. Tailored-Controls Difference Profiles

In the second method, we match each paired galaxy to a set of 20 control galaxies of a similar stellar mass and redshift, as mentioned in Section 2.3. We take the individual averaged profiles of the selected control galaxies and take the median value of the profiles at each radius bin giving us a stacked profile of control galaxies. The error of the profile is the standard error of the mean of the individual profiles. We also take into account the error associated with the random selection of controls using a bootstrapping method to quantify the variation in the values of the stacked profiles over 1000 iterations. We estimate that this random selection contributes an error of about 0.1 dex.

Once the stacked profile of the control galaxies is made, we take the difference between the paired galaxy’s profile and the stacked control profile in log space. We do this process for each of the 220 paired galaxies.

The difference profiles at this stage have now been controlled for stellar mass, redshift, and galaxy size. We can freely stack these profiles to study the global properties of the merging galaxies. We take all of the $\Delta\log(\text{sSFR})$ profiles and take the median value within each radius bin to create a stacked profile of the individual difference profiles. We split the individual difference profiles into five evenly spaced stellar mass bins between $\log(M/M_\odot) = 9.0$ –11.5 and stack the difference profiles within each mass bin. This gives us five difference profiles covering five different stellar mass ranges shown in the left hand panel of Figure 3.

The difference profiles are shown to be centrally enhanced by $\sim 0.25 \pm 0.1$ dex. The enhancement falls to zero around $1.0 R_{\text{eff}}$ and is slightly suppressed in their disks by ~ 0.1 dex. While all of the other paired galaxies show a sSFR suppression in their disks, the profiles of the pairs in the stellar mass range of $\log(M/M_\odot) = 9.0$ –9.5 show an enhancement in their disks of ~ 0.1 – 0.2 ± 0.1 dex.

The profiles are largely consistent with the results of the previous mass binning method. The profiles of this method have less scatter, which may be due to how the paired are more closely matched in stellar mass and redshift. This method

also has the added benefit that the profiles can be separated by other parameters instead of just the stellar mass.

We can also inspect the star formation enhancement within individual paired galaxies. We extract the $\Delta\log(\text{sSFR})$ from the individual difference profiles. We extract the value of $\Delta\log(\text{sSFR})$ from the inner $0.5 R/R_{\text{eff}}$ of the paired galaxies and take the mean value of the $\Delta\log(\text{sSFR})$ within the aperture. We depict the central $\Delta\log(\text{sSFR})$ as a function of stellar mass in the right hand panel of Figure 3. The mass independence can be clearly seen here as the paired galaxies within each stellar mass bin fall within 0.1 dex of the 0.2 dex $\Delta\log(\text{sSFR})$ median.

Between the two different pair-control comparison methods, we conclude that the $\Delta\log(\text{sSFR})$ is independent of the paired galaxies' stellar mass. Low mass galaxies are having their sSFR enhanced at a similar rate as high mass galaxies.

4.2. Future Mass Growth

The higher rate of star formation within interacting galaxies will mean that the stellar mass in the interacting galaxies will grow at a faster rate compared to secular mass growth. The question is whether or not the advanced mass growth rate is significant.

To estimate the enhanced mass growth rate due to merger induced star formation, we will assume a constant star formation rate over a given timescale, τ . A specific star formation rate is already a representation of the fractional mass growth every year, so the merger induced mass growth is simply;

$$\frac{\Delta M_{\text{merger}}}{M_{\text{total}}} = (\text{sSFR}_{\text{pair}} - \text{sSFR}_{\text{control}}) * \tau \quad (9)$$

For this estimation we use a timescale of 1 Gyr, which is the order of the typical timescale between the passage of the first and second pericenters in dark-matter simulations (Boylan-Kolchin et al. 2008). The actual timescale for star formation enhancement is likely shorter; however, we will use the 1 Gyr value as an order of magnitude estimate.

To get the difference in sSFR between the pairs and controls, we fit a gaussian to the difference profiles split by stellar mass. We then add the modeled difference profile to the stacked control profile in the same mass range in log space. We then take the difference between the profiles in linear space and multiply the difference by the 1 Gyr timescale.

We show the mass growth rate profile for each of the four mass ranges in Figure 4. The mass growth rate is largely restricted to the centers of the galaxies and falls to zero around $1.0 R/R_{\text{eff}}$, following profile of the $\Delta\log(\text{sSFR})$. The mass growth rate is strongly dependent on the stellar mass of the paired galaxies.

The fractional mass change is greatest in the less massive galaxies, $\Delta M_{\text{merger}}/M_{\text{total}} = 0.06$ (6%), while it is less substantial in the more massive galaxies, $\Delta M_{\text{merger}}/M_{\text{total}} = 0.005$ (0.5%). This is unsurprising as the $\Delta\log(\text{sSFR})$ profiles were independent of the galaxies' total stellar mass.

In the context of galaxy evolution, where less massive galaxies merge to form larger galaxies, the merger induced star formation contributes a significant amount of stellar mass on top of the stellar mass contribution from the mixing of the two galaxies. Further, the merger induced mass growth is restricted to the centers of the paired galaxies which will result in galaxies which have more prominent bulges in their centers.

5. DISCUSSION

We primarily focused on how the $\Delta\log(\text{sSFR})$ in paired galaxies behaves as a function of stellar mass, but we also explored how the $\Delta\log(\text{sSFR})$ behaves as a function of the mass ratio between the paired galaxies and the projected separation between the paired galaxies. While we see hints that the $\Delta\log(\text{sSFR})$ is affected by the mass ratio and projected separation, these results are less definitive than the $\Delta\log(\text{sSFR})$ profile's lack of dependence on stellar mass.

5.1. Dependency on Mass Ratio

While the total stellar mass between the galaxies has no effect on the merger induced star formation, the mass ratio between the paired galaxies does. The mass ratio here is calculated as the difference between stellar masses contained within the 50% half light radius,

$$\Delta\log(M) = \log(M_{\text{target}}) - \log(M_{\text{comp}}) \quad (10)$$

Where M_{target} is the stellar mass of the MaNGA target galaxy and M_{comp} is the stellar mass of the companion galaxy. Note that we do not use the absolute value of the difference like how the mass ratio is typically used. We do this so that we can independently study how the star formation large central galaxies behave in comparison to small satellite galaxies.

For the outside-IFU sample we use the stellar masses calculated from the r-band elliptical Petrosian apertures from the NSA catalog. Most of the companion galaxies in the inside-IFU sample are not covered by either the NSA catalog or Simard et al. (2011). Because of this, the inside-IFU galaxies are left out of this analysis.

In the Top row of Figure 5 we see that the sSFR enhancement is strongest in galaxies close to a 1:1 mass ratio, $\Delta\log(M) = 0.0$. These galaxies feature a central enhancement of ~ 0.3 dex which is 0.1 dex higher than the average central enhancement of the whole sample. This enhancement falls with wider mass ratio; however, pairs with large differences in stellar mass still feature substantial levels of star formation enhancement, 0.15–0.20 dex at mass ratios of $|\Delta\log(M)| = 1.0$.

We also use the mass ratio to explore how the merger induced tidal torques affect paired galaxies with large mass ratios, $|\Delta\log(M)| > 0.5$. We use the mass ratio to separate the massive central galaxy from satellite galaxies. A paired galaxy with a positive mass ratio is the more massive galaxy of the pair while a paired galaxy with negative mass ratio is the less massive galaxy. We see that the less massive galaxy of a pair show a higher level of sSFR enhancement, by ~ 0.1 dex, with respect to the more massive galaxy. This shows us that satellite galaxies may be more susceptible to tidal interactions than the central galaxies in the pair system.

In previous studies it was noted that the enhancement to the SFR is strongest in galaxy pairs with small mass ratios Ellison et al. (2008). We see the same effect; however, we have also shown that galaxy pairs with large mass ratios still show a substantial level of sSFR enhancement. It has been a convention to split galaxy pairs in major mergers, $|\log(M)| \leq 0.5$, and minor mergers $0.5 < |\log(M)| \leq 1.0$. In this work, we can see that the mass ratio range used in galaxy pair studies could be widened to at least $|\log(M)| \leq 1.5$ so that larger pair samples can be constructed.

Further, within these galaxy pairs with wide mass ratios we see that the less massive galaxy in the pair shows higher levels of sSFR enhancement. This implies that satellite galaxies are

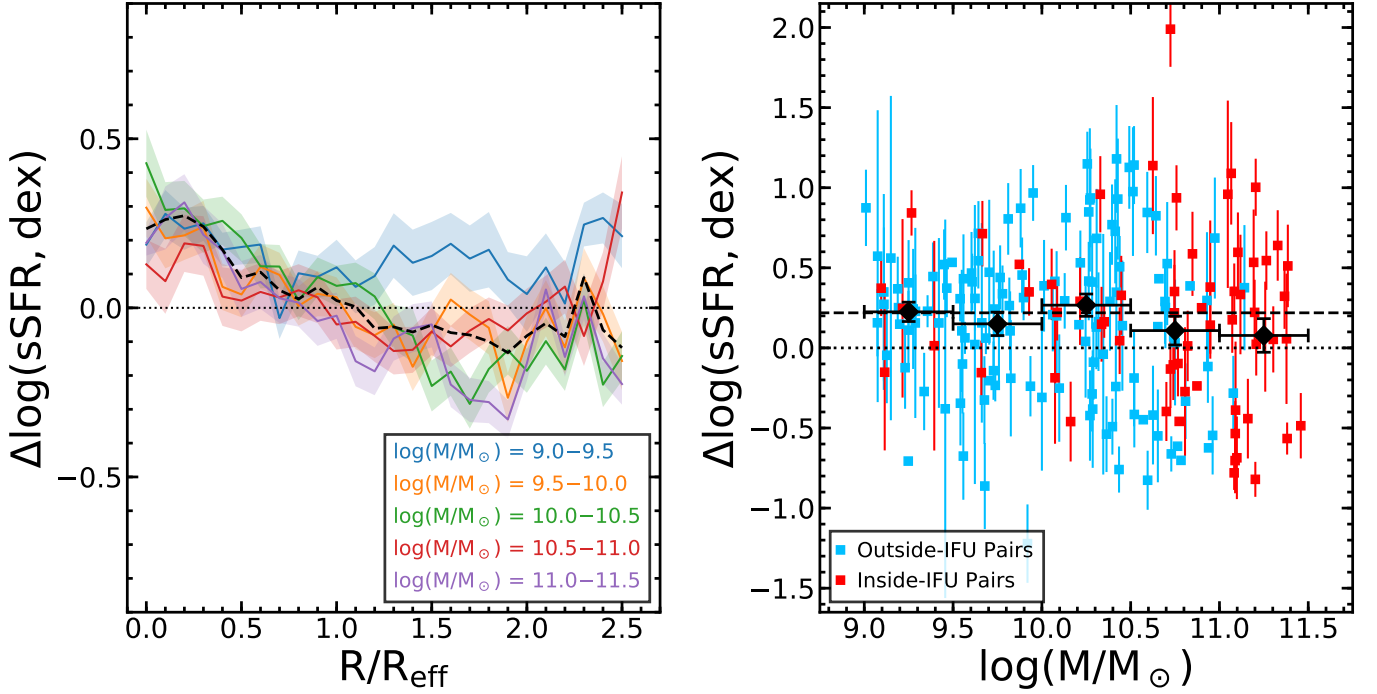


Figure 3. The **Left** panel shows the $\Delta\log(\text{sSFR})$ profiles where the difference profiles are constructed from the difference between the paired galaxy profiles and a set of 20 control galaxies. The profiles are split into four different stellar mass bins and the highlighted region about the profiles represent the standard error of mean of the profile. The black dashed line represents the mean profile between the four difference mass ranges. In the **Right** panel the nuclear $\Delta\log(\text{sSFR})$ are shown. The black squares are the mean values within a stellar mass bin (where the size of the bins are shown in the horizontal error bars). The vertical error bars on the black squares represent the standard deviation within the bin. The horizontal, dashed black line represents the median central enhancement of the pair sample. Galaxies in the outside-IFU (Blue) and inside-IFU (Red) samples are separately depicted.

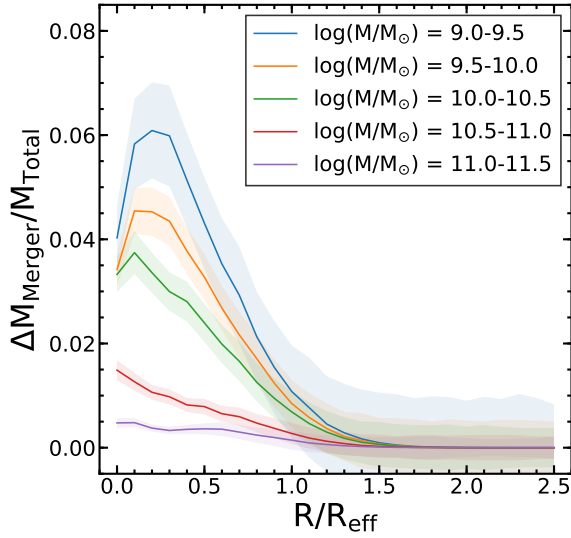


Figure 4. The fractional mass growth due to merger induced star formation over 1 Gyr. The highlighted region represents the propagated standard errors of the mean of the stacked $\log(\text{sSFR})$ profiles.

more affected by interaction induced star formation than the central galaxy.

5.2. Dependency on Projected Separation

We also look at the level of sSFR enhancement as a function of the projected separation between the two galaxies in the bottom row of Figure 5. The $\Delta\log(\text{sSFR})$ profiles show a clear gradient with the projected separation where close galaxy pairs show higher levels of sSFR enhancement com-

pared to pairs with wide separations. The profiles in the galaxy pairs in the separation range, $r_p = 0.0-12.5$, are generally ~ 0.1 dex above the median and the profiles of the galaxy pairs in the separation range, $r_p = 37.5-50.0$, are generally ~ 0.1 dex below the median.

This effect can be seen most clearly in the scatter plot of the central $\Delta\log(\text{sSFR})$ enhancement as a function of projected separation. The level of the enhancement gradually increases with closer separation from 50 kpc to 10 kpc. While $\Delta\log(\text{sSFR})$ falls at higher separations, there is still a substantial level of enhancement between 40 and 50 kpc, $\sim 0.1-0.2$ dex. Within 15 kpc, the $\Delta\log(\text{sSFR})$ enhancement jumps to 0.3–0.4 dex.

The sSFR enhancement increasing with closer projected separations is to be expected and has been shown in previous studies (Li et al. 2008; Ellison et al. 2008; Scudder et al. 2012; Patton et al. 2013). The passage of the first pericenter is predicted to trigger a burst of star formation which falls as the two galaxies separate (Scudder et al. 2012). The sSFR enhancement in our sample persists out to 50 kpc, Patton et al. (2013) showed that the burst of star formation can persist out to 150 kpc.

For galaxies with separations of $r_p < 5$ kpc and $r_p = 30-35$ kpc we find a $\Delta\log(\text{sSFR})$ of zero. For the $r_p < 5$ kpc galaxies, about half of the pairs show signs of tidal features while the others do not. The pairs without the tidal features may be approaching their first pericenter, meaning that the merger-induced gas-inflows have not been triggered yet. The dips in $\Delta\log(\text{sSFR})$ may also just be due to statistical error from the relatively limited sample size.

Between the three parameters, stellar mass, mass ratio, and projected separation, the projected separation has the

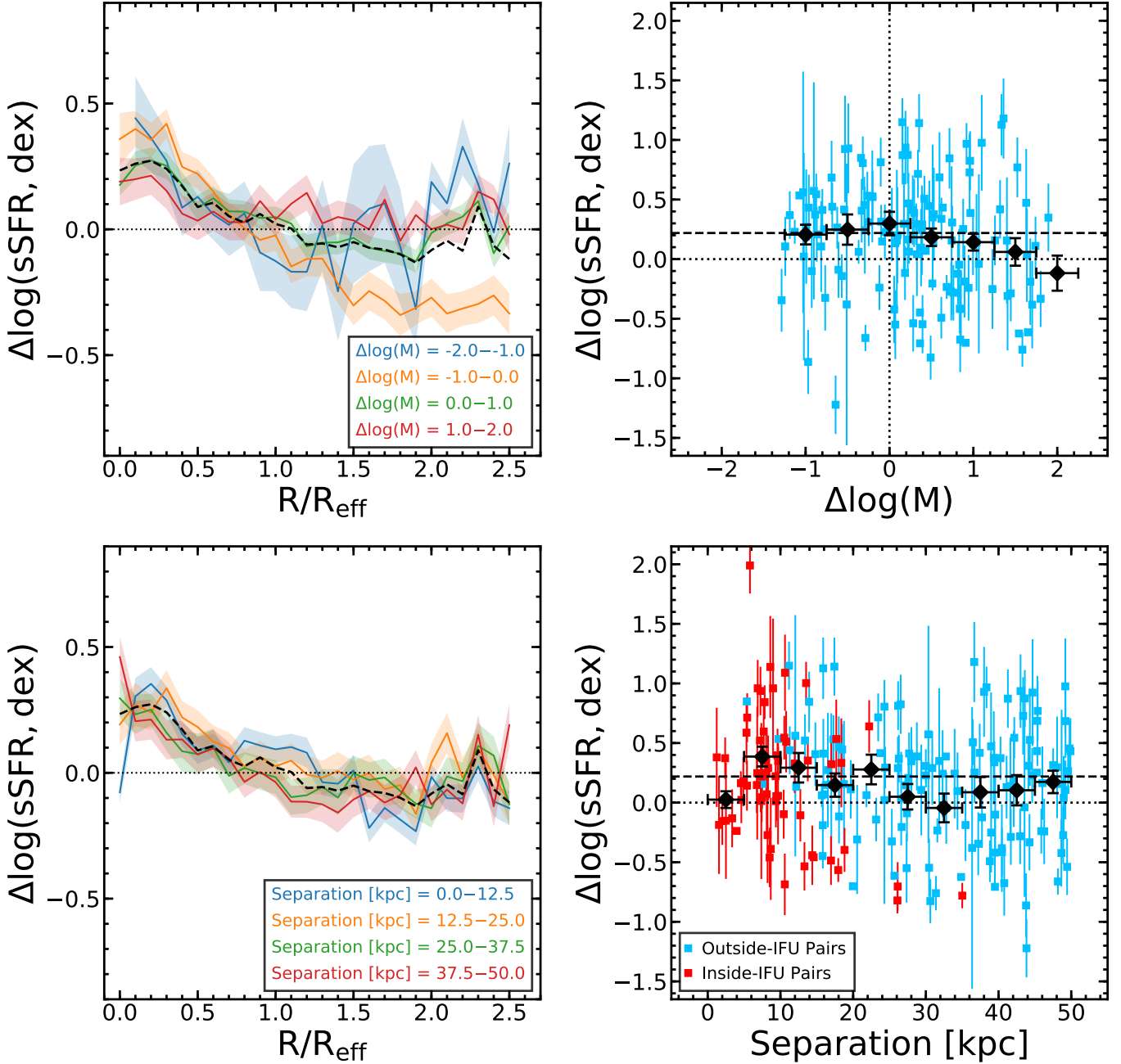


Figure 5. *Use larger bins here.* Same as Figure 3 except the difference profiles are split by mass ratio (**Top row**) and projected separation (**Bottom row**).

strongest effect on the $\Delta\log(\text{sSFR})$ enhancement, followed by the mass ratio. From this we see that the paired galaxies with the strongest enhancements to the central sSFR are paired galaxies with 1:1 mass ratios and which are within 10 kpc.

Between the splitting the $\Delta\log(\text{sSFR})$ profile by stellar mass, mass ratio, and projected separation, we see that the general shape of the profile within the inner $1.0 R/R_{\text{eff}}$ is unchanged. The profile is peaked in the center of the paired galaxies and linearly (in log space) to zero around $1.0 R/R_{\text{eff}}$. Different projected separations and mass ratios may elevate or suppress the level of the enhancement in the profile's center; however, the shape of the profile is largely unchanged.

5.3. Comparison with Previous Works

Li et al. (2008) studied the star formation enhancement over different stellar masses and saw that low mass galaxies experience greater levels of enhancement than high mass galaxies. Galaxies in the mass range $\log(M/M_{\odot}) = 9.72$ were shown to experience an enhancement ~ 0.4 dex higher than galaxies in the mass range $\log(M/M_{\odot}) = 10.6$. In this work, we find that $\Delta\log(\text{sSFR})$ has no dependence on the stellar mass of the galaxy. The galaxies with stellar masses above $\log(M/M_{\odot}) \geq 10.5$ do show slightly lower levels of enhancement compared to the lower mass galaxies; however, the pairs still do not significantly deviate from the average $\Delta\log(\text{sSFR})$ of the sample.

The $\Delta\log(\text{sSFR})$ as a function of projected separation has

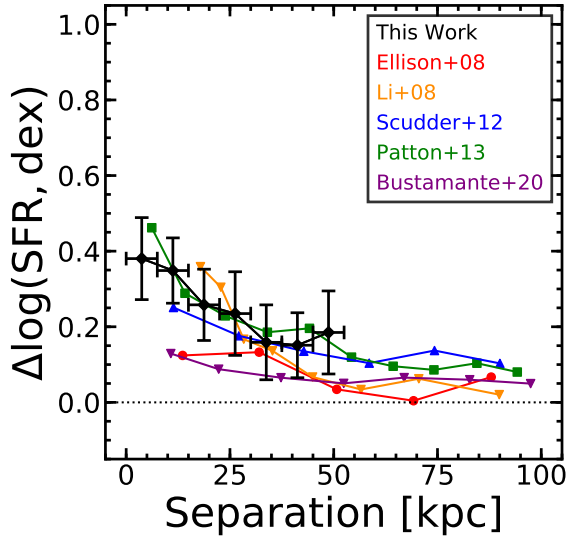


Figure 6. $\Delta\log(\text{SFR})$ over projected separation for this sample (Black), Ellison et al. (2008) (Red), Li et al. (2008) (Orange), Scudder et al. (2012) (Blue), Patton et al. (2013) (Green), and Bustamante et al. (2020) (Purple).

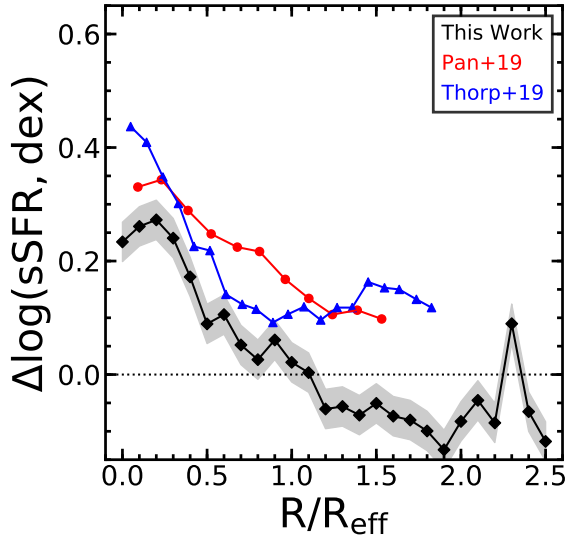


Figure 7. The mean radial profile of $\Delta\log(\text{sSFR})$ between pairs and controls of this work (red) compared against those of Pan et al. (2019) (Blue) and Thorp et al. (2019) (Green).

been studied before with the SDSS survey (Ellison et al. 2008; Li et al. 2008; Patton et al. 2013; Scudder et al. 2012; Bustamante et al. 2020). We compare the mentioned surveys with our own work in Figure 6. We find that our central SFR enhancements are higher than many of the previous studies, the enhancements are ~ 0.15 dex higher than Scudder et al. (2012) and ~ 0.25 dex higher than Ellison et al. (2008) and Bustamante et al. (2020). The central enhancement to sSFR is ~ 0.1 dex higher in our pairs compared to Li et al. (2008) past 35 kpc, but Li et al. (2008) is ~ 0.1 dex higher than our pairs at 20 kpc. Our $\Delta\log(\text{SFR})$ as a function of projected separation is very consistent to what was found in Patton et al. (2013).

Our sample only covers projected separations within 50 kpc, while previous surveys cover out to 100–200 kpc. While our sample covers a smaller separation range, the $\Delta\log(\text{SFR})$ of our sample at 50 kpc is roughly the same level as Scudder et al. (2012) and Patton et al. (2013) at the same projected

separation. This shows that the central SFR enhancement of our sample is consistent with what has been found in previous works.

In Figure 7 we look at the $\Delta\log(\text{sSFR})$ profile as a function of galactocentric radius between our work and previous works. Pan et al. (2019) examines a set of paired galaxies in the MaNGA survey while Thorp et al. (2019) examines a set of post-merger galaxies in the MaNGA survey. The galaxy pairs in Pan et al. (2019) have sSFR enhancements that are roughly ~ 0.1 – 0.2 dex higher than our sample. The $\Delta\log(\text{sSFR})$ profile is closest to our work at the centers of the galaxy pairs but then begins to diverge further out into the galaxies’ disks. Our paired galaxies show a slight sSFR suppression in their disks past $1.0 R/R_{\text{eff}}$ while the galaxy pairs in Pan et al. (2019) show an enhancement of ~ 0.15 dex to the sSFR in their disks.

The post-merger galaxies in Thorp et al. (2019) have sSFR enhancements between ~ 0.05 – 0.2 higher than our sample of paired galaxies. The centers of the post-merger galaxies feature a greater level of sSFR enhancement in their centers compared to our galaxies pairs of $\Delta\log(\text{sSFR}) = 0.5$, roughly 0.25 dex higher than our paired galaxies. The post-merger galaxies also feature enhanced sSFR in their disk past $1.0 R/R_{\text{eff}}$ of ~ 0.1 – 0.15 dex. The heightened level of sSFR enhancement in the centers of post-merger galaxies in comparison to merging galaxies seems consistent with the idea that a second and greater burst of star formation occurs as the two merging galaxies coalesce into a single galaxy.

Barrera-Ballesteros et al. (2015) studied the sSFR in paired galaxies in the CALIFA IFS survey using the $\text{H}\alpha$ equivalent width ($\text{EW}(\text{H}\alpha)$) as a proxy. They compared the sSFR in the centers of paired galaxies to their disks by varying the aperture diameters from which they extract the $\text{EW}(\text{H}\alpha)$. They found that the sSFR is enhanced in the centers of the paired galaxies but suppressed in their disks. Between our work, Barrera-Ballesteros et al. (2015), and Pan et al. (2019) there are three different scenarios; one where the disks of the paired galaxies see no change to the sSFR, one where the sSFR is suppressed, and one where the sSFR is enhanced.

Looking at individual difference profiles in our work, we see that there is a large variance in the $\Delta\log(\text{sSFR})$ where there are as many paired galaxies with sSFR enhancement in their disks as there are paired galaxies with suppression in their disks. The end result is that the average of the $\Delta\log(\text{sSFR})$ is zero. The merger induced effects seen in the disks of the paired galaxies may be more complicated than what is seen in the centers of the galaxies.

6. SUMMARY AND CONCLUSION

In this work, we studied a sample of paired, star forming galaxies using the MaNGA IFS survey. We have made a sample of 220 star forming galaxy pairs along with a set of 1891 star forming control galaxies. We compared the radial profiles of the paired galaxies with those of the controls galaxies and found the following;

1. The sSFR is centrally enhanced within the inner $1.0 R/R_{\text{eff}}$ by 0.25 ± 0.1 dex.
2. The disks of the paired galaxies, $R/R_{\text{eff}} = 1.0$ – 2.5 , show limited suppression in their disks, ~ 0.1 dex.
3. The sSFR enhancement in paired galaxies is independent of the total stellar mass of the paired galaxies,

Massive galaxies experience the same level of sSFR enhancement as low mass galaxies.

4. The sSFR enhancement in paired galaxies does depend on the ratio between the masses of the paired galaxies. Galaxies with small mass ratios, $|\Delta\log(M)| = 0.0\text{--}0.5$, see an enhancement to sSFR of 0.3 ± 0.1 dex, which is 0.1 dex higher than the average paired galaxies in the survey. We have also showed that paired galaxies with large mass ratios, $|\Delta\log(M)| = 1.0$, still show substantial levels of enhancement to the sSFR, 0.15–0.2 dex.
5. Further, in paired galaxies with large mass ratios, the less massive galaxies experience higher levels of sSFR enhancement than the more massive galaxy in the pair, ~ 0.1 dex.
6. The sSFR enhancement increases with closer projected separations where galaxies within 5–15 kpc have their sSFR enhanced by $0.3\text{--}0.4 \pm 0.1$ dex.
7. When splitting the sSFR enhancement profiles by stellar mass, mass ratio, or projected separation, the general shape of the enhancement profile is unchanged. While lower projected separations or mass ratios closer to 1:1 may raise the level of the central enhancement to the profiles, the profiles still linearly fall to zero enhancement around $1.0 R/R_{\text{eff}}$ with roughly the same gradient.
8. The paired galaxies will experience a fractional mass growth rate of 1–6%, where the growth in mass in largest in low mass galaxies.

J.S. and H.F. acknowledge support from the National Science Foundation (NSF) grant AST-1614326. Funding for the Sloan Digital Sky Survey IV has been provided by the Alfred P. Sloan Foundation, the U.S. Department of Energy Office of Science, and the Participating Institutions. SDSS acknowledges support and resources from the Center for High-Performance Computing at the University of Utah. The SDSS web site is www.sdss.org.

SDSS is managed by the Astrophysical Research Consortium for the Participating Institutions of the SDSS Collaboration including the Brazilian Participation Group, the Carnegie Institution for Science, Carnegie Mellon University, the Chilean Participation Group, the French Participation Group, Harvard-Smithsonian Center for Astrophysics, Instituto de Astrofísica de Canarias, The Johns Hopkins University, Kavli Institute for the Physics and Mathematics of the Universe (IPMU) / University of Tokyo, the Korean Participation Group, Lawrence Berkeley National Laboratory, Leibniz Institut für Astrophysik Potsdam (AIP), Max-Planck-Institut für Astronomie (MPIA Heidelberg), Max-Planck-Institut für Astrophysik (MPA Garching), Max-Planck-Institut für Extraterrestrische Physik (MPE), National Astronomical Observatories of China, New Mexico State University, New York

University, University of Notre Dame, Observatório Nacional / MCTI, The Ohio State University, Pennsylvania State University, Shanghai Astronomical Observatory, United Kingdom Participation Group, Universidad Nacional Autónoma de México, University of Arizona, University of Colorado Boulder, University of Oxford, University of Portsmouth, University of Utah, University of Virginia, University of Washington, University of Wisconsin, Vanderbilt University, and Yale University.

REFERENCES

- Abolfathi, B., Aguado, D. S., Aguilar, G., et al. 2018, *ApJS*, 235, 42
 Ahumada, R., Prieto, C. A., Almeida, A., et al. 2020, *ApJS*, 249, 3
 Baldwin, J. A., Phillips, M. M., & Terlevich, R. 1981, *PASP*, 93, 5
 Barnes, J. E., & Hernquist, L. 1996, *ApJ*, 471, 115
 Barnes, J. E., & Hernquist, L. E. 1991, *ApJL*, 370, L65
 Barrera-Ballesteros, J. K., Sánchez, S. F., García-Lorenzo, B., et al. 2015, *A&A*, 579, A45
 Barton, E. J., Geller, M. J., & Kenyon, S. J. 2000, *ApJ*, 530, 660
 Boylan-Kolchin, M., Ma, C.-P., & Quataert, E. 2008, *MNRAS*, 383, 93
 Bundy, K., Bershad, M. A., Law, D. R., et al. 2015, *ApJ*, 798, 7
 Bustamante, S., Ellison, S. L., Patton, D. R., & Sparre, M. 2020, *MNRAS*, 494, 3469
 Capelo, P. R., Dotti, M., Volonteri, M., et al. 2017, *MNRAS*, 469, 4437
 Cardelli, J. A., Clayton, G. C., & Mathis, J. S. 1989, *ApJ*, 345, 245
 Cid Fernandes, R., Stasińska, G., Mateus, A., & Vale Asari, N. 2011, *MNRAS*, 413, 1687
 Drory, N., MacDonald, N., Bershad, M. A., et al. 2015, *AJ*, 149, 77
 Ellison, S. L., Patton, D. R., Simard, L., & McConnachie, A. W. 2008, *AJ*, 135, 1877
 Fu, H., Steffen, J. L., Gross, A. C., et al. 2018, *ApJ*, 856, 93
 Giovanelli, R., Haynes, M. P., Salzer, J. J., et al. 1994, *AJ*, 107, 2036
 Kartaltepe, J. S., Sanders, D. B., Scoville, N. Z., et al. 2007, *ApJS*, 172, 320
 Lambas, D. G., Tissera, P. B., Alonso, M. S., & Coldwell, G. 2003, *MNRAS*, 346, 1189
 Law, D. R., Yan, R., Bershad, M. A., et al. 2015, *AJ*, 150, 19
 Li, C., Kauffmann, G., Heckman, T. M., Jing, Y. P., & White, S. D. M. 2008, *MNRAS*, 385, 1903
 Lin, L., Koo, D. C., Weiner, B. J., et al. 2007, *ApJL*, 660, L51
 Mendel, J. T., Simard, L., Palmer, M., Ellison, S. L., & Patton, D. R. 2014, *ApJS*, 210, 3
 Mihos, J. C., & Hernquist, L. 1996, *ApJ*, 464, 641
 Murphy, E. J., Condon, J. J., Schinnerer, E., et al. 2011, *ApJ*, 737, 67
 Noeske, K. G., Weiner, B. J., Faber, S. M., et al. 2007, *ApJL*, 660, L43
 O'Donnell, J. E. 1994, *ApJ*, 422, 158
 Pan, H.-A., Lin, L., Hsieh, B.-C., et al. 2019, *ApJ*, 881, 119
 Patton, D. R., Torrey, P., Ellison, S. L., Mendel, J. T., & Scudder, J. M. 2013, *MNRAS*, 433, L59
 Perez, J., Michel-Dansac, L., & Tissera, P. B. 2011, *MNRAS*, 417, 580
 Pillepich, A., Nelson, D., Hernquist, L., et al. 2018, *MNRAS*, 475, 648
 Robotham, A. S. G., Driver, S. P., Davies, L. J. M., et al. 2014, *MNRAS*, 444, 3986
 Rodriguez-Gomez, V., Pillepich, A., Sales, L. V., et al. 2016, *MNRAS*, 458, 2371
 Rupke, D. S. N., Kewley, L. J., & Chien, L. H. 2010, *ApJ*, 723, 1255
 Scudder, J. M., Ellison, S. L., Torrey, P., Patton, D. R., & Mendel, J. T. 2012, *MNRAS*, 426, 549
 Simard, L., Mendel, J. T., Patton, D. R., Ellison, S. L., & McConnachie, A. W. 2011, *ApJS*, 196, 11
 Stasińska, G., Vale Asari, N., Cid Fernandes, R., et al. 2008, *MNRAS*, 391, L29
 Thorp, M. D., Ellison, S. L., Simard, L., Sánchez, S. F., & Antonio, B. 2019, *MNRAS*, 482, L55
 Toomre, A., & Toomre, J. 1972, *ApJ*, 178, 623
 Wake, D. A., Bundy, K., Diamond-Stanic, A. M., et al. 2017, *AJ*, 154, 86
 Wong, K. C., Blanton, M. R., Burles, S. M., et al. 2011, *ApJ*, 728, 119
 Woods, D. F., Geller, M. J., & Barton, E. J. 2006, *AJ*, 132, 197
 Xu, C. K., Zhao, Y., Scoville, N., et al. 2012, *ApJ*, 747, 85