

# PHYS6006 Final Report

## Magnetospheric structure associated with high-latitude auroras

J. Plank<sup>1</sup> and R. C. Fear<sup>2</sup>

<sup>1, 2</sup>Department of Physics and Astronomy, University of Southampton

April 2020

### Abstract

*Context:* Only case studies have been done on the magnetospheric structure associated with the formation of high latitude aurora, and the connection of interplanetary magnetic field direction is suspected but has never been quantified.

*Aims:* To produce a statistical survey on the relationship between a northward pointing interplanetary magnetic field and a phenomenon known as transpolar arcs.

*Methods:* Using data from the ESA satellite Cluster, we analysed the ion temperature for many different values of  $Z_{GSM}$ , covering the plasma sheet as well as the magnetotail lobes.

*Results:* We found a direct link to IMF direction based on high temperature events observed in the lobe, suggesting high-latitude aurora form during periods of northward pointing IMF.

<b>Summary</b>	<b>Page</b>
<b>1 INTRODUCTION</b>	<b>4</b>
1.1 Structure of Earth's magnetic field . . . . .	4
1.2 Coordinate Systems . . . . .	4
1.3 The Interplanetary Magnetic Field . . . . .	5
1.4 Reconnection and Aurora . . . . .	5
<b>2 OBSERVATIONS</b>	<b>7</b>
2.1 Limitations of Cluster's orbit . . . . .	7
2.2 High temperatures in the lobe . . . . .	8
2.3 Temperature distribution in the magnetosphere . . . . .	8
2.4 High temperature events . . . . .	8
2.5 Duration of events . . . . .	9
2.6 Interplanetary Magnetic Field . . . . .	10
<b>3 RESULTS</b>	<b>11</b>
3.1 Connection to Transpolar Arcs . . . . .	11
<b>4 DISCUSSION</b>	<b>11</b>
<b>5 CONCLUSION: FUTURE WORK</b>	<b>11</b>

<b>List of Figures</b>	<b>Page</b>
1 Representation of the solar-terrestrial magnetic interaction. The sun is located to the left, producing solar winds and the IMF. Solar winds collide with the terrestrial magnetic field at supersonic speeds, creating a bow shock wave which encases the magnetosphere in a magnetosheath of compressed solar wind [3]. (Image courtesy of NASA) . . . . .	4
2 The GSE (left) and GSM (right) coordinate systems. Image credit: [6]. . . . .	5
3 The Aurora Borealis, photographed in 2015 [22]. . . . .	6
4 Sketch of the Dungey cycle. The IMF is assumed to be coming from the left. On the right hand side of the diagram, magnetotail reconnection is happening. Open field lines are being compressed together, eventually joining to form a closed field line that 'snaps back' towards Earth, the reconnection happens typically at a radius of $\approx 100R_e$ . Heating and acceleration of the plasma occurs as the newly created closed field line rapidly shortens in length. The magnetic flux will then flow around to the dayside of the planet (LHS of the sketch) where it gets pushed up against the solar wind. Once again, this pressure eventually reaches a point where the field lines of the IMF and the terrestrial field merge, creating an open line with its foot on the polar cap. This line gets pushed back over to the night-side of the planet where magnetotail reconnection can happen once again (Sketch from [7]). . . . .	6

5	An example of a transpolar arc observed by Fear and Milan [8]. Dayglow is visible at the top of the image, the auroral oval is also visible as a circle in the centre of the image. The transpolar arc is the thin red band that stretches from the night side of the planet (bottom) into the polar cap region. An observer standing underneath a transpolar arc and looking up would see something that looks just like the normal aurora. The only significant difference is they are standing at a latitude that under normal circumstances would be far too high to see any aurora. . . . .	7
6	Plot of Cluster's orbit throughout September 2005 (solid), along with model results for the magnetopause and open-closed boundary on 15/09/2005 (dotted) [9, 18]. The grey shaded region is of interest as it contains open field lines. If a reconnection event was detected while the spacecraft was in this region then it could indicate a transpolar arc. Plotting is done in the GSM coordinate system, where +X points to the sun and +Z points along the magnetic dipole axis. . . . .	7
7	The temperature distribution in the magnetosphere. Measurements from Cluster orbits in the months May-December during the years 2002-2010. Each pixel is $1R_e^2$ in area, and represents the mean ion temperature when cluster was in that position, blue for lower temperatures and red for hotter. The plasma sheet is visible as a high temperature region to the right of Earth. . . . .	8
8	Box plot of the temperature distribution as a function of $Z_{GSM}$ . Each box covers a width of $1R_e$ , starting at $[0, 1) R_e$ then $[1, 2) R_e$ etc until $[11, 12] R_e$ . Yellow lines represent the median, each box is the interquartile range of each bin, and the "whiskers" are maximum and minimum values, up to $3 \times IQR$ . Outliers plotted as black 'x'. The box centred at $5.5R_e$ (for the interval $[5, 6) R_e$ ) suggests that temperatures above $\approx 20MK$ are good candidates for high temperature events. Data is from Jul-Oct 2002-2010 for $X_{GSM} < 0 R_e$ . . . . .	9
9	Histograms showing the number of events with a certain duration (in minutes). Each panel shows a different cutoff point for $ Z_{GSM} $ . Very long duration events are consistent with passing through the plasma sheet where having a temperature above $20MK$ for over 10 hours is possible. Therefore as the Z cutoff is increased we would expect those long duration events to become less common as measurements outside of the plasma sheet are excluded. What is left is 732 (see $> 6R_e$ panel) high temperature events occurring outside the region where they are expected. . . . .	10

**1. INTRODUCTION** Earth and the Sun are connected by more than just gravity, magnetic interactions between the two bodies are the cause of some of the most dramatic and complex phenomena on Earth. The most well known of these is the aurora (Borealis in the northern hemisphere and Australis in the southern). More commonly known as the northern and southern lights, these light shows come as a result of charged particles flowing along the Sun's magnetic field lines (known as solar wind) and making their way down to Earth in a zone known as the auroral oval [5]. The Sun's violent nature is also transmitted to Earth along its magnetic field, known as the 'Interplanetary magnetic field' (IMF). This can result in satellite damage, radiation hazards to astronauts and airline passengers, telecommunications problems, and outages of power and electronics systems [5].

The modern theory of the formation of the aurora was first proposed by Kristian Birkeland in 1903 [4]. He said that auroras are produced when solar wind encounters the geomagnetic field. The result of this is the formation of the plasmasphere in the equatorial plane of Earth's magnetic field [1].

**1.1. Structure of Earth's magnetic field** Figure 1 is a representation of the solar-terrestrial environment. The solar wind, travelling with the interplanetary magnetic field from the sun [21], is coming from the left. It collides with the terrestrial magnetic field at supersonic speeds, creating a bow shock wave in front of the magnetosphere - the region where magnetic field lines connect to Earth at both ends - the boundary of the magnetosphere is the magnetopause. Supersonic solar winds compress at this boundary creating a magnetosheath between the magnetopause and the bow shock [3].

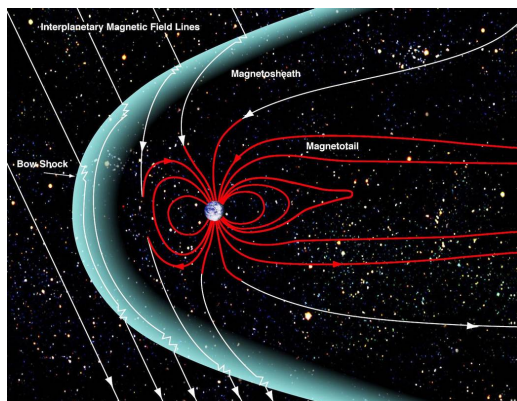


Figure 1: Representation of the solar-terrestrial magnetic interaction. The sun is located to the left, producing solar winds and the IMF. Solar winds collide with the terrestrial magnetic field at supersonic speeds, creating a bow shock wave which encases the magnetosphere in a magnetosheath of compressed solar wind [3]. (Image courtesy of NASA)

**1.2. Coordinate Systems** Geocentric Solar Ecliptic (GSE) and Geocentric Solar Magnetic (GSM) are two similar coordinate systems used when studying the magnetosphere. Both have the Earth as the origin, and +X pointing towards the sun. In GSE, +Z is pointed perpendicularly upwards from the plane of Earth's orbit around the Sun. In GSM, Z is the projection of Earth's magnetic dipole axis onto the plane perpendicular to X, with positive pointing north.

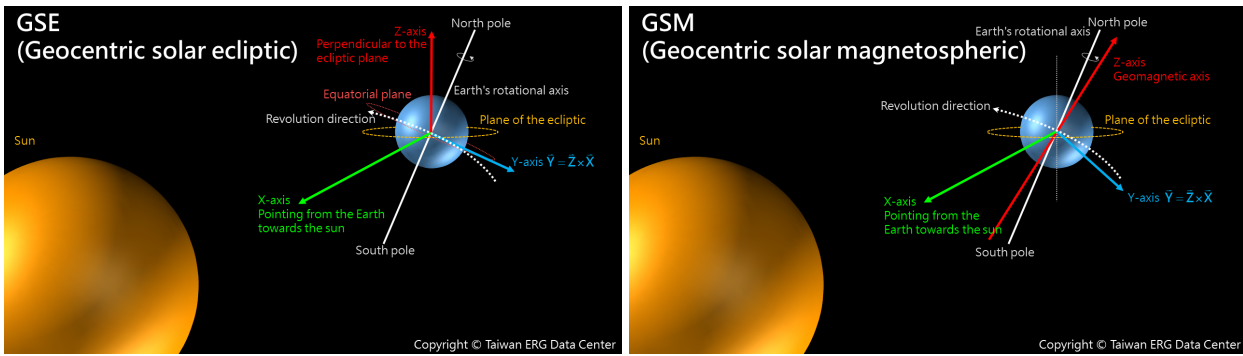


Figure 2: The GSE (left) and GSM (right) coordinate systems. Image credit: [6].

**1.3. The Interplanetary Magnetic Field** Magnetic field lines are said to be “frozen in” to the solar wind plasma, the IMF is carried into the interplanetary region of space by solar wind from the sun. The mechanism for this is known as Alfvén’s theorem, first proposed in 1942 [2].

#### DERIVE FROZEN-IN!

Because of the rotation of the sun, solar wind (along with IMF) follows a spiral pattern similar to that produced by droplets of water as a wet tennis ball is thrown into the air with rapid spin. These field lines are ‘open’, i.e. they have an origin, but instead of returning to a corresponding region at the opposite pole, extend indefinitely into space [20].

At the Sun’s magnetic equator, field lines originating from the north and south hemispheres run parallel to each other but are oppositely directed - creating a thin current sheet known as the “Heliospheric current sheet“. This sheet is twisted and warped due to the difference in rotational and magnetic axes and a quadrupole moment in the sun’s magnetic field - it also has a spiral shape of the same form as described above [2, 17].

Since the orbital plane of the Earth is located almost along the rotation axis of the sun, we experience regular shifts in the direction of flow of the IMF because the heliospheric current sheet is sometimes above and sometimes below the planet.

When the IMF has a southward component, the northward pointing terrestrial field lines are allowed to merge in a process known as reconnection. However, when the IMF is northward, the structure of the magnetosphere is not well understood [9].

**1.4. Reconnection and Aurora** The auroral lights, pictured in Fig.3, are a result of charged particles from the sun being accelerated towards the ionosphere - Earth’s upper atmosphere - where they collide with gas molecules, each molecule causing a different colour to be emitted. E.g. Green (the most common colour) is caused by collisions with atomic oxygen (557.7nm) [12, 3]. Intense aurora will have an emission rate of several million Rayleigh ( $1R = 10^6 \text{ photons/cm}^2 \text{ s}$ ) and most often appear as east-west aligned bands known as auroral arcs [3].



Figure 3: The Aurora Borealis, photographed in 2015 [22].

Reconnection is a process that allows magnetic fields from separate domains to become joined to one another. When IMF is southward, reconnection occurs in a smooth cycle known as the Dungey cycle [7] which is explained in Fig.4. Aurora is observed at latitudes of approximately  $70^\circ$  because that is the boundary between the closed lines and the open ones at the polar cap (lobe).

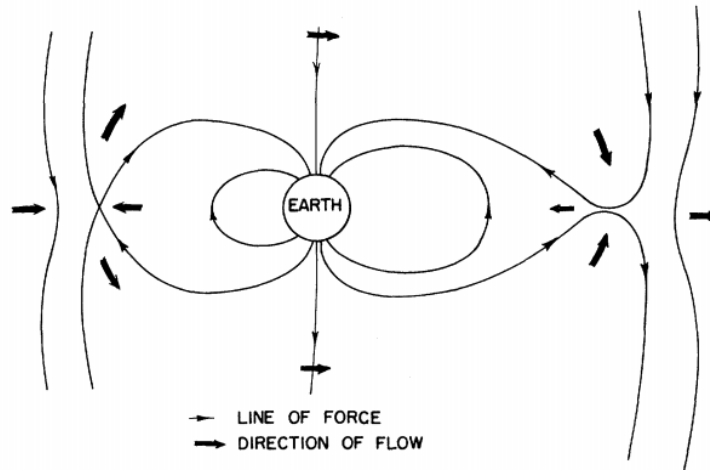


Figure 4: Sketch of the Dungey cycle. The IMF is assumed to be coming from the left. On the right hand side of the diagram, magnetotail reconnection is happening. Open field lines are being compressed together, eventually joining to form a closed field line that ‘snaps back’ towards Earth, the reconnection happens typically at a radius of  $\approx 100R_e$ . Heating and acceleration of the plasma occurs as the newly created closed field line rapidly shortens in length. The magnetic flux will then flow around to the dayside of the planet (LHS of the sketch) where it gets pushed up against the solar wind. Once again, this pressure eventually reaches a point where the field lines of the IMF and the terrestrial field merge, creating an open line with its foot on the polar cap. This line gets pushed back over to the night-side of the planet where magnetotail reconnection can happen once again (Sketch from [7]).

Aurora at very high latitudes are rare, but can occur in an event known as a transpolar arc [10, 11] (see figure 5). There is much debate on their formation, one theory for their formation was proposed by Milan et al in [16], where magnetic flux in the magnetotail lobes - the open field line region that maps down to the polar cap - reconnects but is trapped in the magnetotail. Transpolar arcs occur

when the IMF is northward, which would explain why flux gets trapped in the magnetotail, there is a blockage created because dayside reconnection cannot occur.

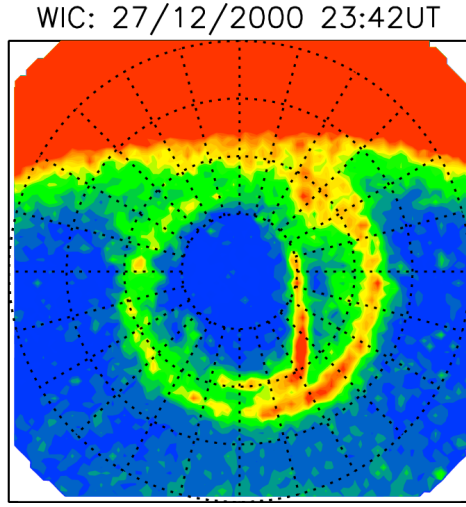


Figure 5: An example of a transpolar arc observed by Fear and Milan [8]. Dayglow is visible at the top of the image, the auroral oval is also visible as a circle in the centre of the image. The transpolar arc is the thin red band that stretches from the night side of the planet (bottom) into the polar cap region. An observer standing underneath a transpolar arc and looking up would see something that looks just like the normal aurora. The only significant difference is they are standing at a latitude that under normal circumstances would be far too high to see any aurora.

**2. OBSERVATIONS** The majority of observations in this report come from Cluster, a group of four satellites built by ESA and launched in pairs on Soyuz rockets from Baikonur on the 16th July and 9th August 2000. They fly in a 57hr elliptical polar orbit and are arranged in a tetrahedron formation with a separation of between a few hundred and a few thousand km.

**2.1. Limitations of Cluster's orbit** The orbit of cluster varies throughout the year. For significant portions of the year the satellite is well outside of the lobes, therefore the best months to extract data are July-October every year, where the spacecraft spends the majority of its time in the tail, and does not enter the magnetosheath.

Figure 6 shows Cluster's orbit through the month of September 2005 as an example, it does not cross the predicted magnetopause boundary (outer dotted parabola). Cluster will also cross through the plasmasphere in every orbit. This leaves the grey shaded region as the area of interest.

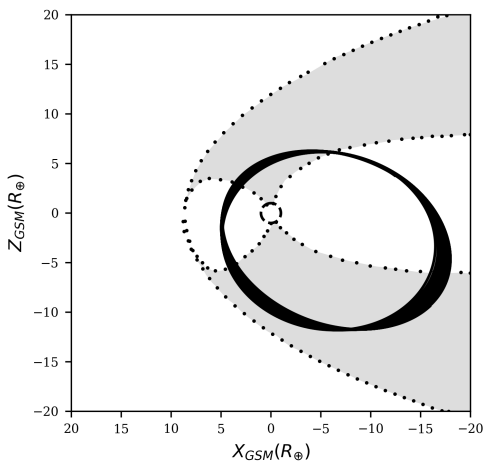


Figure 6: Plot of Cluster's orbit throughout September 2005 (solid), along with model results for the magnetopause and open-closed boundary on 15/09/2005 (dotted) [9, 18]. The grey shaded region is of interest as it contains open field lines. If a reconnection event was detected while the spacecraft was in this region then it could indicate a transpolar arc. Plotting is done in the GSM coordinate system, where +X points to the sun and +Z points along the magnetic dipole axis.

**2.2. High temperatures in the lobe** Occasionally when Cluster is in the lobe, it will detect periods of uncharacteristically high temperature. There is some debate on the cause of this; Shi *et al* [19] attribute it to solar wind penetrating into the lobe, whereas Fear *et al* [9] suggests that this cannot be the case due to the presence of a double loss cone.

A double loss cone occurs when there is a magnetic mirror at both ends of the field line. This would imply that the observed field line is closed, since the Earth's dipole field has the property where magnetic field strength increases as you travel along a field line to either pole. This would cause an ion to slow down and reverse direction at the poles, i.e. it bounces back and forth from pole to pole.

**2.3. Temperature distribution in the magnetosphere** Figure 7 shows the mean distribution of temperatures in the magnetosphere, during the period May to December for years 2002, 2010. We can see the general shape of the plasma sheet, a hot region to the right of the Earth where the normal aurora- generating reconnection happens.

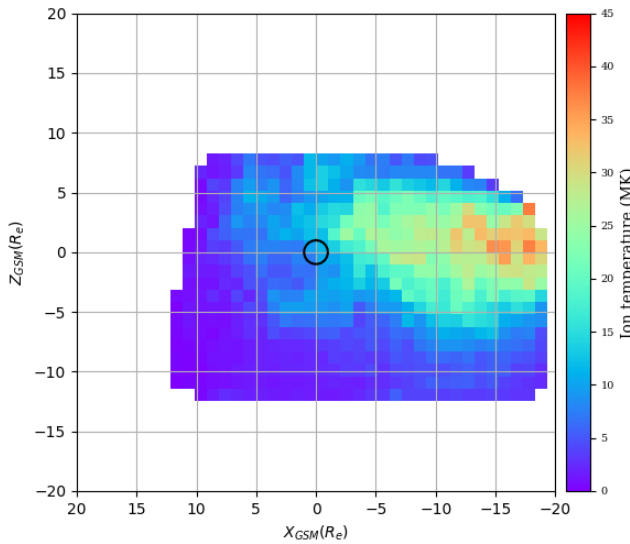


Figure 7: The temperature distribution in the magnetosphere. Measurements from Cluster orbits in the months May-December during the years 2002-2010. Each pixel is  $1R_e^2$  in area, and represents the mean ion temperature when cluster was in that position, blue for lower temperatures and red for hotter. The plasma sheet is visible as a high temperature region to the right of Earth.

**2.4. High temperature events** What constitutes a high temperature? That depends on where in the magnetosphere you are looking. The most easily quantifiable way to separate the plasma sheet from the lobe is to look at the Z coordinate in the Geocentric Solar Magnetospheric (GSM) coordinate system [15].



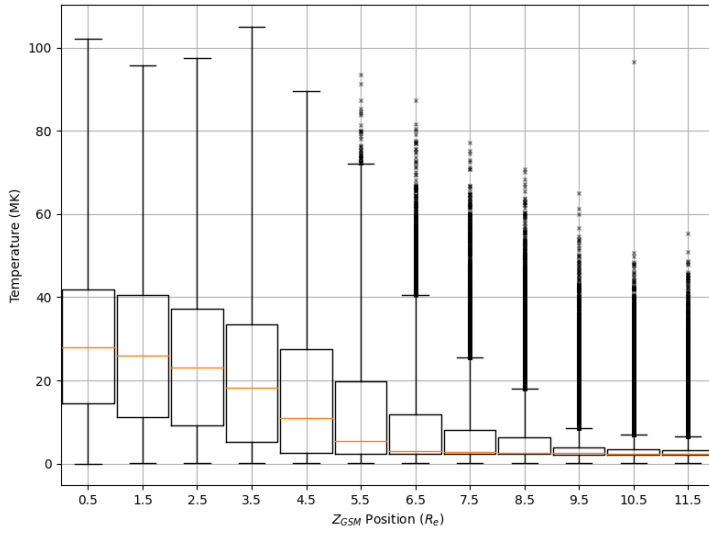


Figure 8: Box plot of the temperature distribution as a function of  $Z_{GSM}$ . Each box covers a width of  $1R_e$ , starting at  $[0, 1) R_e$  then  $[1, 2) R_e$  etc until  $[11, 12] R_e$ . Yellow lines represent the median, each box is the interquartile range of each bin, and the “whiskers” are maximum and minimum values, up to  $3 \times IQR$ . Outliers plotted as black ‘x’. The box centred at  $5.5R_e$  (for the interval  $[5, 6) R_e$ ) suggests that temperatures above  $\approx 20MK$  are good candidates for high temperature events. Data is from Jul-Oct 2002-2010 for  $X_{GSM} < 0 R_e$ .

We can see from Figure 7 that the plasma sheet ends at approximately  $Z_{GSM} = \pm 6R_e$ . Giving an exact position for the edge of the plasma sheet is difficult as it is dynamic and in this project we were using measurements from many months and years.

In figure 8 we created a box plot of measurements from July-October 2002-2010 (for  $X_{GSM} < 0$ , i.e. excluding points sunwards of Earth). Taking the estimation from above that the plasma sheet ends at  $\approx 6R_e$ , we can see from the box centred on  $5.5R_e$  that temperatures exceeding  $\approx 20MK$  can reasonably be considered uncharacteristically high.

**2.5. Duration of events** The event analysed by Fear *et al* [9] on 15-09-2005 lasted for approximately 2 hours. In figure 9 we analyse the length of each of our events detected in the same time period as above (Jul-Oct 2002-2010), plotting the result as a histogram. Repeating this for various cutoff points based on  $|Z_{GSM}|$  gives a view of how the length of each event changes over time.

count	732	25%	00:15:00
mean	01:16:18	50%	00:25:00
std	03:38:36	75%	01:00:00
min	00:10:00	max	1 day 12:40:00

Table 1: Statistical properties for the duration of events occurring at  $|Z_{GSM}| > 6R_e$ .

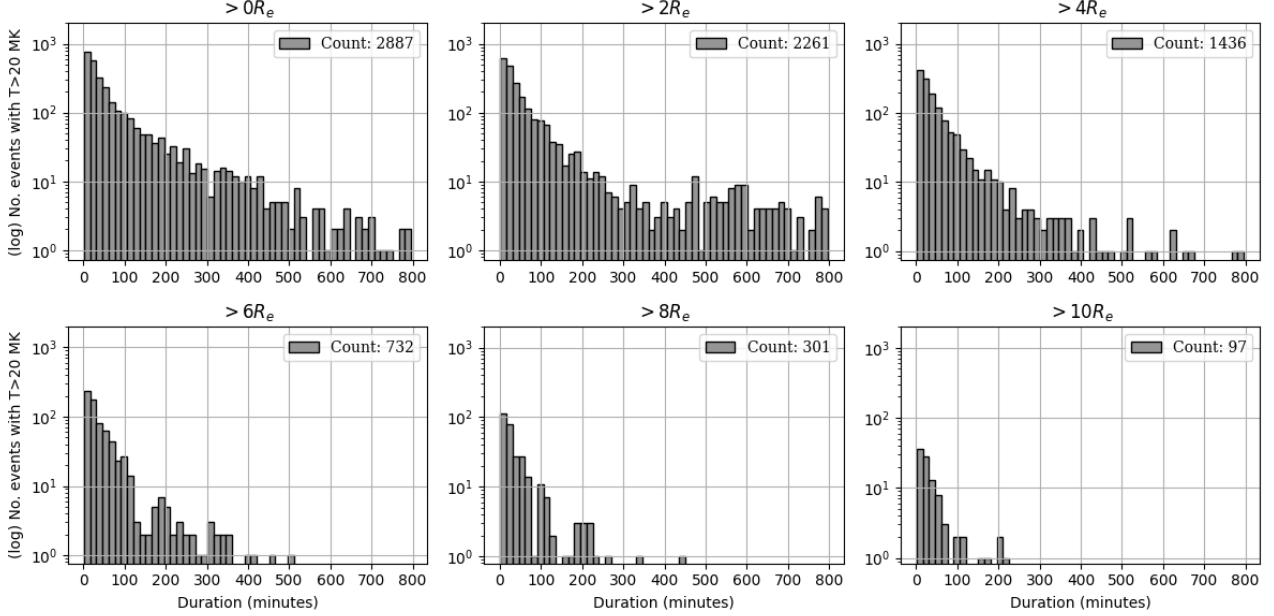


Figure 9: Histograms showing the number of events with a certain duration (in minutes). Each panel shows a different cutoff point for  $|Z_{GSM}|$ . Very long duration events are consistent with passing through the plasma sheet where having a temperature above 20MK for over 10 hours is possible. Therefore as the Z cutoff is increased we would expect those long duration events to become less common as measurements outside of the plasma sheet are excluded. What is left is 732 (see  $> 6R_e$  panel) high temperature events occurring outside the region where they are expected.

We see that when the plasma sheet is excluded by excluding results where  $|Z_{GSM}| < 6R_e$ , there are 732 events left. This works out to be just under one per day ( $732/984 \approx 0.75 \text{ events/day}$ ). The properties of this distribution are shown in table 1. This shows that these events are comparable in length to the 15-09-2005 event.

Many short duration events are likely to be noise. Although cluster makes a full rotation approximately every three seconds and can therefore provide measurements with similar resolution, the choice was made to average the data into five minute bins, which would match with the predicted geometric position data from [13], and the IMF components parameters obtained from NASA/GSFC's OMNI data set through OMNIWeb [14].

**2.6. Interplanetary Magnetic Field** The IMF data consists of three measurements from OMNI,  $B_x, B_y, B_z$ , all in the GSM coordinate system.

$B_z$  represents the IMF polarity, positive northwards and negative southwards. For the time range that we have been considering, the average bz was:

$$\overline{B_z} = 0.004 \pm 3.275 \text{ nT} \quad (1)$$

Therefore it was effectively 50% positive and 50% negative.

### 3. RESULTS

**3.1. Connection to Transpolar Arcs** By comparing the time an event happened along with its duration, it is possible to compare temperatures in the lobe with images of the auroral oval. As shown in section 2.5, there were 732 observed events with a temperature of over 20MK. These can be filtered further to exclude all events with a duration less than 1.5 hours.

This is to align with observations from SSUSI, an instrument carried on Defense Meteorological Satellite Program (DMSP) satellites that has been taking images of the auroral oval every 100 minutes since 2005.

### 4. DISCUSSION HI

### 5. CONCLUSION: HI

### FUTURE WORK

### References

- [1] H. Alfvén. *Cosmical Electrodynamics*. Oxford University Press, 1950.
- [2] H. Alfvén. “Existence of Electromagnetic- Hydrodynamic Waves”. In: *Nature* 150.3805 (1942), pp. 405–406. DOI: 10.1038/150405d0.
- [3] Wolfgang Baumjohann and R. Treumann. *Basic Space Plasma Physics - Revised Edition*. Jan. 2012. DOI: 10.1142/P850.
- [4] K. Birkeland. *The Norwegian Aurora Polaris Expedition 1902-1903: (1st, 2d, Sect.) Birkeland, K. On The Cause Of Magnetic Storms And The Origin Of Terrestrial Magnetism*. Creative Media Partners, LLC, 2018. ISBN: 9780343378370. URL: <https://books.google.co.uk/books?id=x7icvQEACAAJ>.
- [5] Volker Bothmer and Ioannis A. Daglis. *Space Weather: Physics and Effects*. Springer Science and Business Media, 2007.
- [6] *Coordinate Systems*. URL: [http://erg.ncku.edu.tw/erg\\_new2/erg\\_en/DataDoc/Coordi.php](http://erg.ncku.edu.tw/erg_new2/erg_en/DataDoc/Coordi.php).
- [7] J. W. Dungey. “Interplanetary Magnetic Field and the Auroral Zones”. In: *Phys. Rev. Lett.* 6 (2 Jan. 1961), pp. 47–48. DOI: 10.1103/PhysRevLett.6.47. URL: <https://link.aps.org/doi/10.1103/PhysRevLett.6.47>.
- [8] R C Fear and Stephen E. Milan. “The IMF dependence of the local time of transpolar arcs: Implications for formation mechanism”. In: 2012.
- [9] R. C. Fear et al. “Direct observation of closed magnetic flux trapped in the high-latitude magnetosphere”. In: *Science* 346.6216 (2014), pp. 1506–1510. ISSN: 0036-8075. DOI: 10.1126/science.1257377. eprint: <https://science.sciencemag.org/content/346/6216/1506.full.pdf>. URL: <https://science.sciencemag.org/content/346/6216/1506>.

- [10] L. A. Frank et al. "Polar views of the Earth's aurora with Dynamics Explorer". In: *Geophysical Research Letters* 9.9 (1982), pp. 1001–1004. DOI: 10.1029/GL009i009p01001. eprint: <https://agupubs.onlinelibrary.wiley.com/doi/pdf/10.1029/GL009i009p01001>. URL: <https://agupubs.onlinelibrary.wiley.com/doi/abs/10.1029/GL009i009p01001>.
- [11] L. A. Frank et al. "The theta aurora". In: *Journal of Geophysical Research: Space Physics* 91.A3 (1986), pp. 3177–3224. DOI: 10.1029/JA091iA03p03177. eprint: <https://agupubs.onlinelibrary.wiley.com/doi/pdf/10.1029/JA091iA03p03177>. URL: <https://agupubs.onlinelibrary.wiley.com/doi/abs/10.1029/JA091iA03p03177>.
- [12] Andi Hollier. *Southern Lights near the Great Australian Bight*. URL: <https://earthobservatory.nasa.gov/images/91062/southern-lights-near-the-great-australian-bight>.
- [13] *JSOC Catalogues Selection*. URL: [http://cdhf5.bnsc.rl.ac.uk/pub-query/output\\_form?JSOC=1](http://cdhf5.bnsc.rl.ac.uk/pub-query/output_form?JSOC=1).
- [14] J. H. King and N. E. Papitashvili. "Solar wind spatial scales in and comparisons of hourly Wind and ACE plasma and magnetic field data". In: *Journal of Geophysical Research: Space Physics* 110.A2 (2005). DOI: 10.1029/2004JA010649. eprint: <https://agupubs.onlinelibrary.wiley.com/doi/pdf/10.1029/2004JA010649>. URL: <https://agupubs.onlinelibrary.wiley.com/doi/abs/10.1029/2004JA010649>.
- [15] Stefan Maus. *Coordinate Transformations*. July 2006. URL: <https://www.ngdc.noaa.gov/geomag/gsm2geo.shtml>.
- [16] S. E. Milan, B. Hubert, and A. Grocott. "Formation and motion of a transpolar arc in response to dayside and nightside reconnection". In: *Journal of Geophysical Research: Space Physics* 110.A1 (2005). DOI: 10.1029/2004JA010835. eprint: <https://agupubs.onlinelibrary.wiley.com/doi/pdf/10.1029/2004JA010835>. URL: <https://agupubs.onlinelibrary.wiley.com/doi/abs/10.1029/2004JA010835>.
- [17] E. N. Parker. "Dynamics of the Interplanetary Gas and Magnetic Fields." In: 128 (Nov. 1958), p. 664. DOI: 10.1086/146579.
- [18] E. J. Rolfe and Kaldeich-Schürmann Brigitte. *Third international conference on substorms (ICS-3): Versailles, France, 12-17 may 1996*. European Space Agency, 1996.
- [19] Q. Q. Shi et al. "Solar wind entry into the high-latitude terrestrial magnetosphere during geomagnetically quiet times". In: *Nature Communications* 4.1 (Feb. 2013), p. 1466. ISSN: 2041-1723. DOI: 10.1038/ncomms2476. URL: <https://doi.org/10.1038/ncomms2476>.
- [20] Edward J. Smith, Bruce T. Tsurutani, and Ronald L. Rosenberg. "Observations of the interplanetary sector structure up to heliographic latitudes of 16°: Pioneer 11". In: *Journal of Geophysical Research: Space Physics* 83.A2 (1978), pp. 717–724. DOI: 10.1029/JA083iA02p00717. eprint: <https://agupubs.onlinelibrary.wiley.com/doi/pdf/>

10.1029/JA083iA02p00717. URL: <https://agupubs.onlinelibrary.wiley.com/doi/abs/10.1029/JA083iA02p00717>.

- [21] Leif Svalgaard and Edward W. Cliver. “Heliospheric magnetic field 1835-2009”. In: *Journal of Geophysical Research: Space Physics* 115.A9 (Sept. 2010), n/a–n/a. ISSN: 0148-0227. DOI: 10.1029/2009ja015069. URL: <http://dx.doi.org/10.1029/2009JA015069>.
- [22] torino071. *Lofoten*<sub>2</sub>0150215<sub>5</sub>465. Feb. 2015. URL: <https://www.flickr.com/photos/ohrensessel/16398299949/in/photostream/>.