

Rotation effects on the Lyman- α line morphology in distant galaxies

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ABSTRACT

Rotation is present in the gas kinematics of galaxies up to the highest redshifts. In this paper we present for the first time radiative transfer calculations that show the impact of rotation on the morphology of the Lyman α line. To this end we construct simplified models where a galaxy is modeled as an homogeneous sphere composed as an homogeneous mixture of dust and hydrogen at a constant temperature. These spheres have a solid-body rotation with linear velocities at the surface in the range $0 - 300 \text{ km s}^{-1}$. We consider radiation sources both in the center of the rotating cloud and also homogeneously distributed around the sphere. We find that higher rotational velocities increase the width of each peak in the outgoing line profile while it also increases the amount of Lyman alpha photons escaping in the line center. This trends makes that for high rotational velocities and large Hydrogen optical depths the double peak of the line tends to be erased and be replaced by a single peak the lines center. This is more pronounced for radiation sources homogeneously distributed. Concerning the escape fraction we find that rotation does not have any effect, provided that all the sources are centrally emitted. However in the case of homogeneously emitted sources we measure an increase of about a factor of 2 in the escape fraction for higher rotational velocity values. Our work shows clearly that gas rotation has a non negligible impact on the shape of the Lyman α line.

Key words: galaxies: high-redshift - galaxies: star formation - line: formation

1 INTRODUCTION

Due to the resonant nature of the Lyman alpha line, gas kinematics play an important role shaping its morphology...

There is an extensive literature studying the influence of outflow/inflow configurations in the shape of the outgoing Lyman-alpha line...

In this paper we study for the first time the impact of rotation on the morphology of the Lyman α line. To isolate the effects of rotation we focus on a simple system: the gas distribution is spherical, with homogeneous density and the gas rotates as a solid body. We base our work on two independent Monte Carlo based radiative transfer codes CLARA (Forero-Romero et al. 2011) and XX (Dijkstra & Kramer 2012).

This paper is structured as follows...

In this paper we express a photon's frequency in terms of the dimensionless variable $x \equiv (\nu - \nu_a)/\Delta\nu_\alpha$, where $\nu_\alpha = 2.46 \times 10^{15} \text{ Hz}$ is the Ly α resonance frequency, $\Delta_\alpha \equiv \nu_\alpha \sqrt{2kT/m_p c^2} \equiv \nu_a v_{\text{th}}$ is the doppler broadening of the line which depends on the neutral gas temperature T

scattering the radiation or equivalently the thermal velocity v_{th} of the atoms.

2 MODELING BULK GAS ROTATION

Describing the kinematics of gas rotation in all generality is a complex task. There is great variation in the shape of the rotation curve as observed in HI emission as a function of the distance to the galaxy center. However there are two features that are observed very often. First, in the central region the velocity increases proportional to the radius following the behaviour in a body with solid rotation. Second, beyond a certain radius the rotation curve tends to flatten. Furthermore, a thorough observational account of gas rotation in the redshifts of interest for the study of LAEs ($z > 1.0$) is still missing.

An ab-initio description of realistic rotation curves in simulations depends on having access to the dynamic evolution the mass components in the galaxy: stars, gas and dark matter. Such level of realism is extremely complex to achieve,

| Physical Parameter (units) | Symbol | Values |
|--------------------------------|------------|---|
| Velocity (km s ⁻¹) | V_{\max} | 0, 50, 100, 200, 300 |
| Hydrogen Optical Depth | τ_H | 10 ⁵ , 10 ⁶ , 10 ⁷ |
| Dust Optical Depth | τ_a | 0,1 |
| Photons Distributions | | Central, Homogeneous |

Table 1. List of the physical parameters that define the spherical models we have simulated using Monte Carlo calculations. For each parameter we vary the values in the range listed in the last column. Taking into account all the possible combinations we end up with 60 different models.

specially of one wants to get a systematic description based on statistics of simulated objects.

Following the tradition of studies of Lyman α emitting systems, we implement a model with a simplified geometry and gas distribution. We assume that the gas is homogeneously distributed in a sphere that rotates as a solid body with constant angular velocity. This simple model will contain only one parameter: the linear velocity at the sphere's surface, V_{\max} .

2.1 Detailed Implementation of Rotation

In the MonteCarlo code we define a cartesian coordinate system to define the position of each photon. The origin of this system coincides with the center of the sphere and the rotation axis is defined to be z -axis. With this choice, the components of the gas bulk velocity field, $\vec{v} = v_x\hat{i} + v_y\hat{j} + v_z\hat{k}$, can be written as

$$v_x = -\frac{y}{R}V_{\max}, \quad (1a)$$

$$v_y = \frac{x}{R}V_{\max}, \quad (1b)$$

$$v_z = 0, \quad (1c)$$

where R is the radius of the sphere and V_{\max} is the linear velocity at the sphere's surface. The minus/plus sign in the x/y -component of the velocity indicates the direction of rotation. In this case we take the angular velocity in the same direction as the \hat{k} unit vector. With these definitions we can write the angular velocity as $\omega = V_{\max}/R$.

In contrast with spherically symmetric models (static, outflow, inflow) the rotation now defines a preferred direction in the problem. In Section 3 we quantify this difference by varying the line of sight of a mock observer with respect to the rotation axis. The results are parameterized by the polar angle θ as defined by the dot product $\cos\theta = \hat{u} \cdot \hat{k}$.

2.2 Grid of Simulated Galaxies

In the Monte Carlo calculations we follow the propagation of $N_\gamma = 10^5$ numerical photons through different spherical galaxies, each one varying in at least one of the following parameters: the maximum rotational velocity V_{\max} , hydrogen optical depth τ_H , dust optical depth τ_a and the initial distribution of photons with respect to the gas. There are in total 60 models with the input parameters summarized in Table 1.

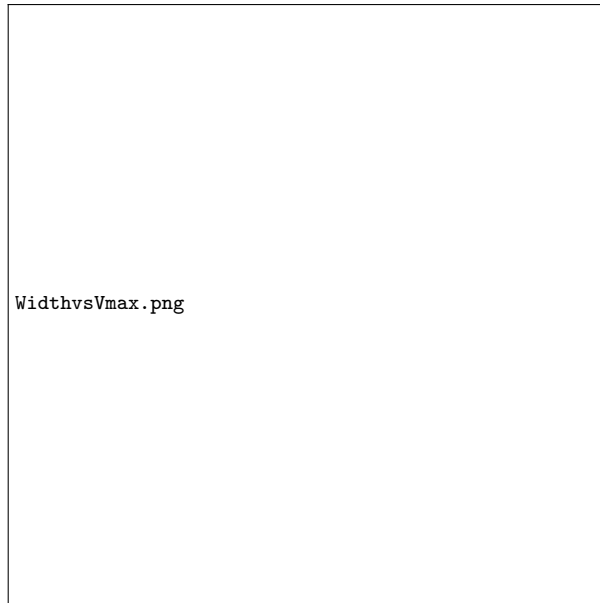


Figure 2. Width of the Lyman-alpha line for all the models.

3 RESULTS

The central result of this paper is summarized in Fig. 1 that shows clearly the considerable impact of rotation on the morphology of the emergent Ly α line. Both panels in the Figure focus on the results for $\tau_H = 10^7$, showing that the influence of rotation is present both when the photons are either homogeneously or centrally initialized over the gas volume.

These results for these outgoing spectra are constructed by taking into account all outgoing photons regardless of the direction of propagation. In subsection ?? we report different statistics aiming at quantifying the changes in the observed spectra for observers with different viewing angles.

In the following subsections we quantify the trends observed in Fig. 1 and Fig. ?? as a function of the maximum rotation velocity V_{\max} and the position of the observer with respect to the rotation angle, $\mu = \cos\theta$. All the results in this section will be expressed in terms of the adimensional frequency x .

First we quantify the line by its full width at half maximum (FWHM) and the peak positions. In order to interpret these results we measure how the rotation affects the average number of scatterings for each Lyman-alpha photon in the simulation. This helps us to introduce the next subsection on the escape fraction in our models including dust. We conclude the section by estimating the expected line flux for top hat filters at a fixed center and varying width.

3.1 Line width and peak maxima

The first quantitative conclusion of the effect of rotation in the Lyman alpha line is that the double peaks in the line tend to broaden until they reach a single broad emission peak. This is most evident in the case of Lyman-alpha sources homogeneously distributed in the gas distributions (Fig. 1 Right panel).

To quantify the line broadening we measure a modified

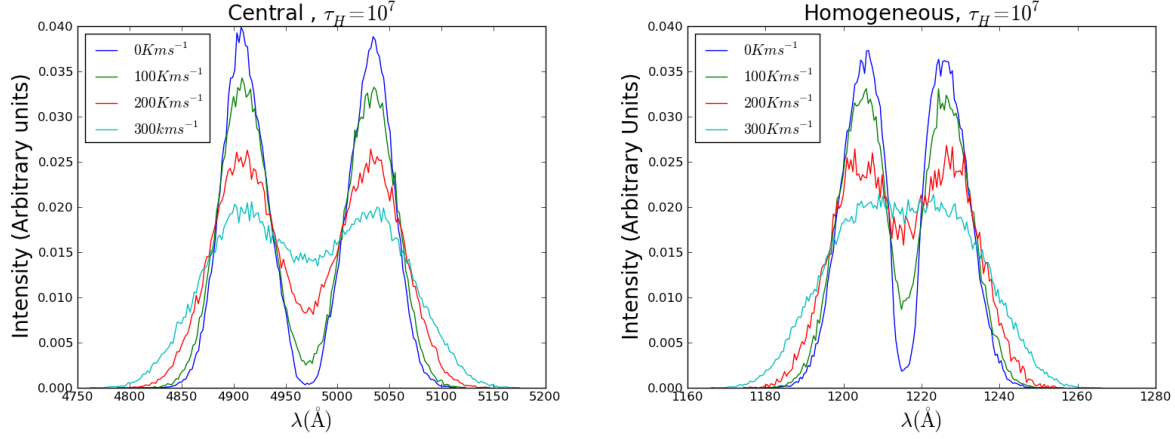


Figure 1. Shape of the Lyman α line for different velocities. The left (right) panel shows the central (homogeneous) photon distribution. All photons were taken into account regardless of their outgoing direction of propagation.

version of the full width at half maximum (FWHM) for half of the line, $W_{1/2}$. It means that in the case of double peaked emission, $W_{1/2}$ corresponds to the width of one of the peaks, while in the extreme case when the line is converted into a single peak, $W_{1/2}$ corresponds to half of the full width.

This definition allows us to quantify the line width both in the cases of double and single peak emission. Furthermore it has the advantage that this line width should have a direct observational correspondence to the observed line feature once the Inter-Galactic Medium (IGM) effects are taken into account, which have the central effect of strongly reducing the intensity of the blue peak of the line.

Fig. 2 summarizes our findings for $W_{1/2}$ as a function of V_{\max} . The line width increases with the rotational velocity of the gas cloud. This increase can be of a factor of 2 – 3 with respect to the width with respect to the static case. This trend is conserved at all optical depths regardless of the initial source distribution.

This result includes all the outgoing photons, regardless of the position of the observer. In Figure ?? we take into account the different positions of the observer in the measurement of the half-width $W_{1/2}$. From this we conclude that observers with a line of sight perpendicular to the axis of rotation (i.e. edge-on in the case of spiral galaxy) tend to measure larger line widths than observers aligned with the rotation axis (i.e. face-on). The influence of the observer position on the line width, amounting always less than 15% of a difference with respect to the result that takes into account all the outgoing photons with the same weight regardless of the relative observer position.

The second feature in the line that we use to quantify the effect of rotation is the position of the line maxima. These provide information on the wavelength of the majority of the outgoing photons after they interact with the neutral hydrogen atoms in the gas cloud. If most of the photon escapes with a low number of scattering, its outgoing frequency will be close to its initial frequency, that is in the center of the line. On the contrary if the number of scatterings is large for the average photon, its outgoing frequency will be far from the line center. Such reasoning can be made more quantitative to understand the dependence of the peak



Figure 3. Position of the maxima in the outgoing spectra for different Rotational velocities, (up) Central Distribution, (Down) Homogeneous Distribution.

maxima as a function of the hydrogen optical depth in the cloud [citation needed].

In Fig. 3 we present the position of the maxima as a function of V_{\max} . For the photons emitted in the line center we do not find any variation in the position of the maxima in the range of explored parameter space (dash lines). However in the homogeneous case we can see how the maxima goes to $x_m = 0$, meaning that the double peak is converted into a single peak.

This transition to a single peak line occurs for the systems where it becomes easier for a bulk of the photons to escape with the lowest number of scatterings possible. This can explain how the single peak stage can be achieved in the homogeneous source distribution where there is a fraction of the photons inside a photosphere region with $\tau_H \sim 1$ which allows them to escape within one scatter. Increasing

the rotational velocity makes it easier for the photons in this photosphere region to escape.

We also take into account the viewing angle μ for all the optical depths and rotational velocities and the two different photon distribution. But we don't find any dependency meaning that the peaks would be placed in the same place no matter the position of an external observer.

Finally we also report on the effect of the neutral Hydrogen optical depth τ_H on the maxima position x_m Fig. 3. We find that, at fixed rotation velocity, the position of the maxima increases with optical depth as expected from basic theoretical considerations. We compare our results with the expected theoretical scaling for an infinite slab.

3.2 Average Number of Scatterings

Until now we have focus in how rotation affect the morphology of the Lyman alpha line, now we turn to study deeply the possible causes of this effects. As recombination is the main cause of the emission of the Lyman alpha line, it is important to study the relationship between the number of times that a photon is absorbed and re emitted that we simply take as scatterings with it's shift in the wavelength.

As a first approach we study the number of scatterings of all the photons at different velocities and for both photon distributions Figure 4. For the central distribution we found that the average number of scattering does not have a relevant change ($\sim 0.5\%$) with V_{max} , For the Homogeneous distribution we found that as V_{max} increase the average number of scatterings decrease in $\sim 39\%$, it means that rotation makes photons to escape easier from the cloud. The reason why this is not observed in the central distribution is because the amount of gas that every photon has to break through is so large that the average number of scatterings is large too, making the effect of rotation unobservable.

In order to understand the effects of rotation in the morphology of the line we make histograms of $\langle N_{scatt} \rangle$ in function of X_m for both the central Fig. ?? and the homogeneous distribution Fig. 6. For the central distribution and for lower velocities we found that for some value (put the value) of N_{scatt} the majority of photons have a defined value of X_m . While for higher velocities photons tend spread out in X_m (Make position analysis of this photons).

For the homogeneous distribution we found the same effect but also, for highest velocities the majority of the photons near the surface escape with much less scatterings that in the static case. This analysis let us conclude that rotation make photons escape with less scatterings and it spread out the wavelength of the outgoing photons represented by X_m .

3.3 Escape Fraction

Until now we have studied models in absence of dust, but is known (REFERENCE) that there is presence of dust in high redshift galaxies. Our models with dust are treated as is explained in detail in the appendix of (Forero et al 2011).

Of particular interest is to compute the escape fraction of Lyman α photons coming from the most distant galaxies, due to the fact that with the observed intensity of the Lyman alpha line quantities as the LF and SFR can be derived (put some references).

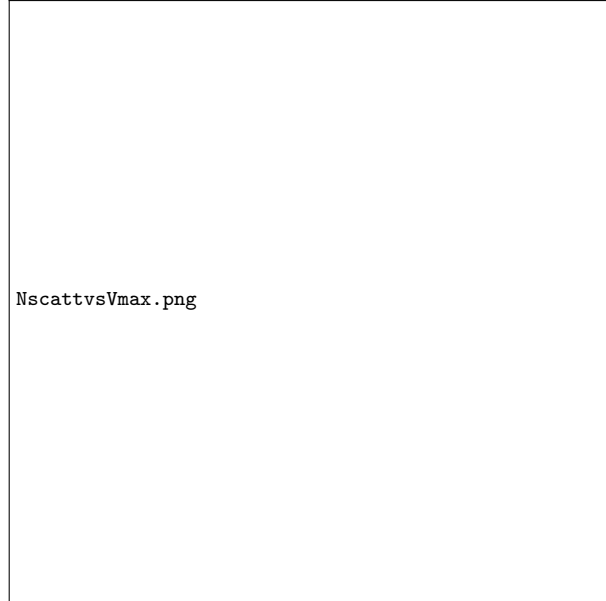


Figure 4. Log of the Average number of scatterings of the outgoing photons as function of the velocity.

Previous studies have shown the correlation of the Escape fraction with galactic mass (Laursen et al 2009, Dayal et al 2010) abundances and the kinematics of dust. In order to study pure rotational effects in the escape fraction we fixed the dust abundance $\tau_A = 0.01$ and the galaxy mass. We compute the escape fraction for the models described in Table 1.

For a realistic model we also take into account the viewing angle, first we fixed the viewing angle at $\theta = 0$ Fig. 7 and then we fixed the velocity in order to see the escape fraction correlation with the viewing angle. Therefore we define the escape fraction as:

$$F_e = \frac{\sum_{NI} \vec{k} \cdot \vec{o}}{\sum_{NF} \vec{k} \cdot \vec{o}} \quad (2)$$

Where NI is the initial number of photons and NF is the final, \vec{k} is the rotation axis direction and \vec{o} the observer direction. With this definition we compute the escape fraction for all of our models, the results are presented in Fig 7

In Figure 7 we found that in the central distribution the escape fraction does not change with velocity while it does in the optical depth (See Verhamme 2006, an argument about this). On the other hand for the homogeneous distribution we found that for higher velocities photons escape easily. The difference between this two results rely in the fact that in the homogeneous distribution photons are emitted closer to the escape surface and this makes this configuration more sensitive to rotation while in the central configuration the escape fraction depends mainly in the amount of gas rather than in rotation.

As a final test we compare our results with the analytical solution of the slab developed by (Neufeld 1990) Fig. 3.5, as the geometry we use is different from the one de-

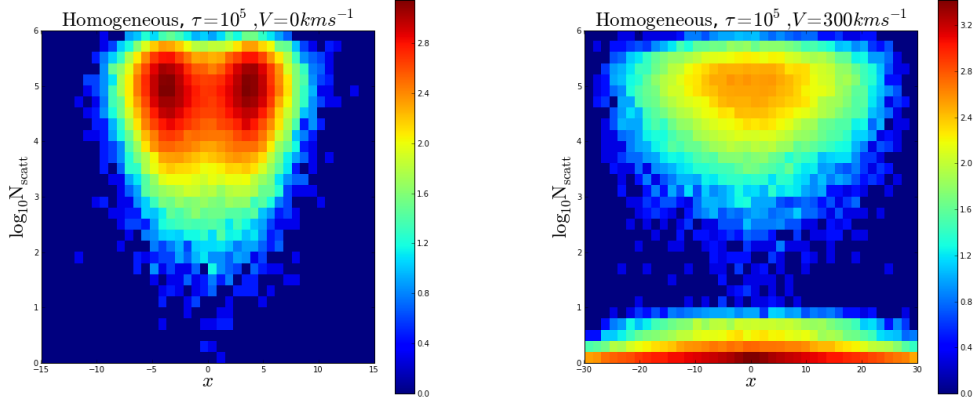


Figure 5. Histogram of N_{scatt} vs x for the homogeneous distribution, left for $V_{max} = 0 \text{ km/s}$, right for $V_{max} = 300 \text{ km/s}$

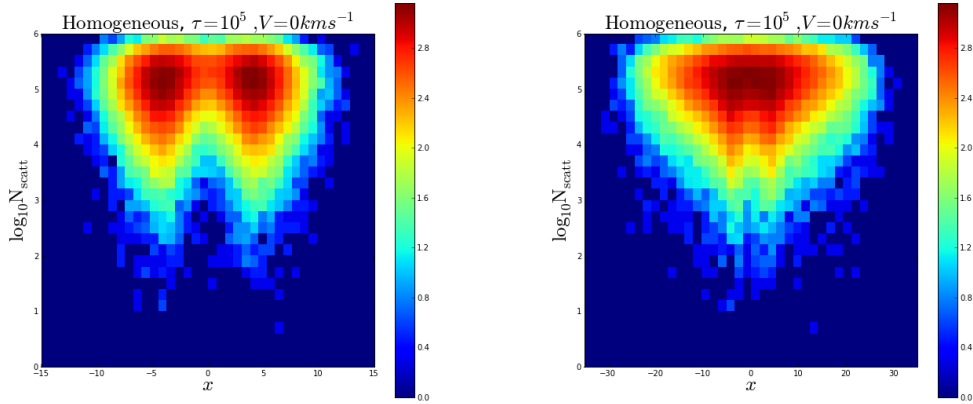


Figure 6. Histogram of N_{scatt} vs x for the Central distribution, left for $V_{max} = 50 \text{ km/s}$, right for $V_{max} = 300 \text{ km/s}$

scribed for Neufeld, we don't expect the same results, in fact we found that for the homogeneous sphere the escape fraction is higher than the slab (see ref, mark). Also we note that for $\tau < 10^6$ the escape fraction does not increase as it will be expected, This is due to the fact that the condition $a\tau_h \gg \tau_d$ is not valid any more.

3.4 The Case of Asymmetric Emission

As we know there is unlikely to find galaxies with radiation sources distributed homogeneously. Most of them are in a clumpy distribution (Laursen et al 2013*) which affected the resulting spectra. In order to study an inhomogeneous distribution we set up a model in which we select certain photons that are placed in a specific place but that are not symmetrically distributed Fig. 9 shows the distributions we set up.

In Fig. 10 we show the resulting spectra of distribution 1 and 2, the first effect we see is the asymmetry of the double peak, in the homogeneous and central distribution we

see double peaks with the same height, while in this case one peak is higher than the other. For $\tau = 10^5$ we found that in distribution 1 the blueshifted peak is higher than the redshifted, and for distribution 2 the redshifted peak is the highest.

Another important fact here is the asymmetry of the spectra with respect to the line center, in particular photons selected in distribution 1 present a blueshift in their spectra while photons selected in distribution 2 present a redshift. This effect becomes stronger as velocity increase 10

3.5 Integrated flux in a narrowband filter

Until now we have shown the main effects of gas and dust rotation in the Ly α line morphology such as the escape fraction, the FWHM, the maxima position, also we see that shape of the outgoing spectra depends on the position of the observer. It is important to see if rotational effects are detectable in observational methods involving the Ly α line.

One of the most used methods to detect high redshift

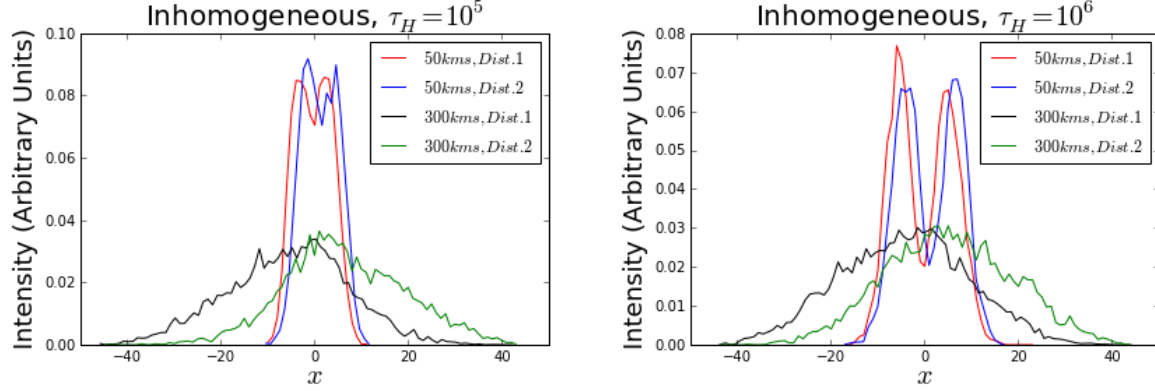


Figure 10. Inhomogeneous model for velocities 50km/s and 300km/s , (left) with $\tau = 10^5$, (right) with $\tau = 10^6$

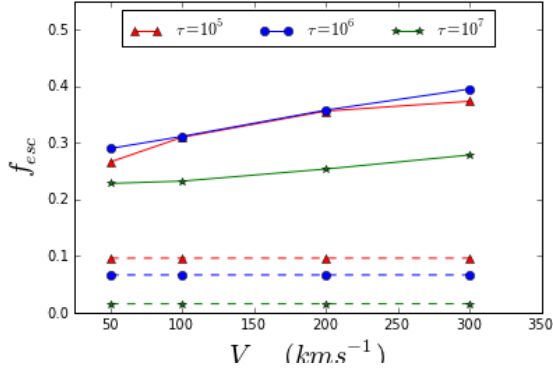


Figure 7. Escape fraction for all the models. continuous lines represent the homogenous models, while dashed lines represent the central models

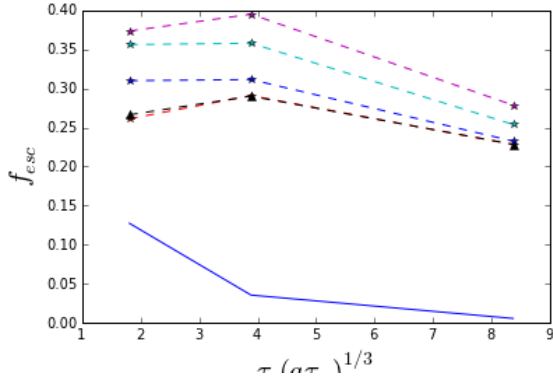


Figure 8. Escape fraction of this work compared with Neufeld analytical solution

galaxies using the Ly α line is using a narrowband selection, we make this analysis based on the results obtained by Steidel (2011) in this work (EXPLAIN a lit of bit more about their work) they used tree narrowband filters for tree different redshifts resumed in Table 2, We want to know how much the integrated flux change due to rotational effects in this NB filters, for the models we simulated with CLARA. In table 2 we present the results of the flux in every

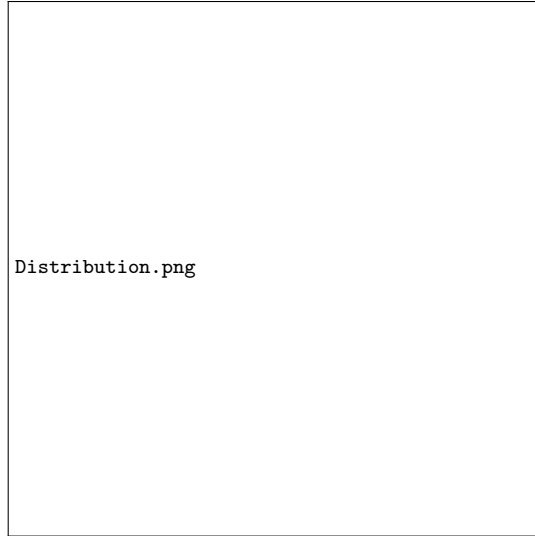


Figure 9. Inhomogeneous distributions of photons, Blue area represents the photons in distribution 1 while red area are the selected photons for distribution 2. The arrow points to the observer position

narrowband filter for the homogeneous model at different velocities and hydrogen optical depth τ_0 . As is well known an increase in the optical depth makes the line peaks separation bigger. In fact for some cases the line width is larger than the NB filter width fig.(put ref fig) for those cases we modify the redshift in the available range in order to make the peak maxima match the NB filter center. For all cases we found that as velocity increase the flux is less.

In the case of the dusty model we found the same trend of the flux with the velocity but in this cases the effect is no that strong. We also found some dependency with the viewing angle.

4 DISCUSSION

... Comparison with Verhamme et al. results on the rotation
... Compare with Kulas et al (Figure 3), Rotation on the lyman alpha line convert double peak profiles into a single one. comments about rotation with inflows and outflows.

| | | | |
|-------|----------------|----------------|----------------|
| Model | SSA22a 4980/80 | HS1549 4667/80 | HS1700 4018/90 |
|-------|----------------|----------------|----------------|

Table 2. Fluxes for tree different narrow band filters.

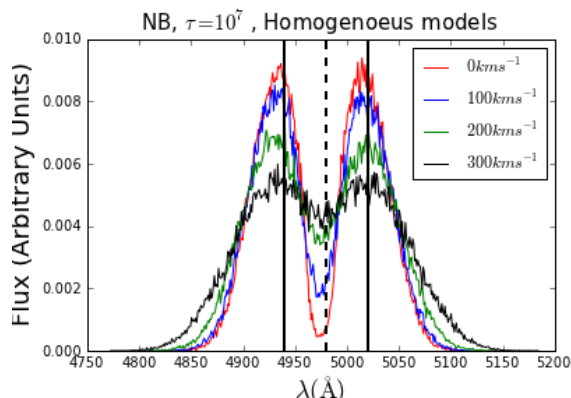


Figure 11. NB filter located at the center of the maxima, with different redshift

... The results derived in this paper have consequences on the interpretation of galaxy observations in the Lyman α line. ..compare steidel et al (2011)

5 CONCLUSIONS

In this paper we have estimated the effects of gas bulk rotation on the emission of the Lyman α line. We based the study on the study of a simplified configuration of an homogeneous sphere rotating as a solid body. We explored a range of models by varying the rotation speed, hydrogen optical depth, dust optical depth and initial distribution of Ly α photons with respect to the gas density. This was implemented in CLARA, a Monte-Carlo radiative transfer code already used to study the Lyman α line.

As first we see how the width of the line changes using a modified FWHM explained in section 3.1, and we found that as gas bulk rotation increase also the width increase in a factor of 2 – 3 in comparison with the static case. We also take into account the influence of the observer viewing angle, we found that observers with a line of sight perpendicular to the axis of rotation measure a 15% larger line width than those aligned with the rotation axis.

As many observational spectra Ly α emission line (Kulas et al) is double peaked, these peaks provide important information concerning gas kinematics and geometry, which can be partially explained with inflows/outflows of gas content. We study the effect of rotation in the position of this peaks, and we find that the position of the maxima does change with rotation for the homogeneous models when the double peak merged into a single peak as velocity increase. This effect is not seen for the central distribution when the double peak remains constant as the velocity increase. We also find that there is no dependency in the observer viewing angle with the maxima position.

Concerning the escape fraction under rotational effects on the Ly α emission line, we found that the escape fraction increase in about 20% – 30% for the homogeneous sphere

model. While rotational effects are negligible for the central models and the escape fraction remains constant. Also the observer viewing angle have no effect in the escape fraction neither for the homogeneous and central models. Complementing this analysis we study the average number of scatterings $\langle N_{scatt} \rangle$ that photons perform before escaping of the cloud taking into account rotational effects. The main result here is for the homogeneous models for which as velocity increase photons escape with about $\sim 39\%$ less scatterings than in the static case.

As an application of these results we compute the integrated flux taking into account the narrow band filters used by (Steidel et al 2011), for our models we found an important decrease up to 40% for the homogeneous models, and up to 22% for the dusty homogeneous models in the flux as velocity increase. Also we calculate at what redshift should the filter be in order to get the maximum flux, and for the tree filters we get values that rely in the filter redshift range. This effects would have a relevant implication at the time to find high redshift galaxies.

This paper illustrates for the first time the main effects of rotation in the morphology of the Ly α emission line, we estimate the range of this effects for simplified models.

ACKNOWLEDGEMENTS

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APPENDIX A: TABLES