

N-body simulations of the Self-Confinement of Viscous Self-Gravitating Narrow Eccentric Planetary Ringlets

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ABSTRACT

N-body simulations are used to illustrate how narrow eccentric planetary ringlets can evolve into a self-confining state.

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1. INTRODUCTION

Narrow eccentric planetary ringlets have properties both interesting and not well understood: sharp edges, sizable eccentricity gradients, and a confinement mechanism that opposes radial spreading due to ring viscosity. Prevailing ringlet confinement mechanisms include: unseen shepherd satellites (reference), periapse pinch (ref), self gravity (ref), and self-confinement (ref). This study uses N-body simulations to show how a viscous narrow self-gravitating ringlet can evolve into a self-confining state.

2. RINGLET CONFINEMENT MECHANISMS

This section will explain the pros and cons of the various ringlet confinement mechanisms, and will then motivate the possibility that ringlets are self confining. That possibility is explored further via numerical simulations using the `epi_int_lite` N-body integrator.

3. EPIINT_LITE

`Epi_int_lite` is a child of the `epi_int` N-body integrator that was used to simulate the outer edge of Saturn’s B ring while it is sculpted by satellite perturbations (Hahn & Spitale 2013). The new code is very similar to its parent but differs in two significant ways: (*i.*) `epi_int_lite` is written in python and recoded for more efficient execution, and (*ii.*) `epi_int_lite` uses a more reliable drift step to handle unperturbed motion around an oblate planet (detailed in Appendix A).

Otherwise `epi_int_lite`’s treatment of ring self-gravity and viscosity are identical to that used by the parent code; see Hahn & Spitale (2013) for additional details. The `epi_int_lite` source code is available

at https://github.com/joehahn/epi_int_lite, and the code’s numerical quality is assessed in Appendix B where the output of several numerical experiments are compared against theoretical expectations.

Calculations by `epi_int_lite` use natural units with gravitation constant $G = 1$, central primary mass $M = 1$, and the ringlet’s inner edge has initial radius $r_0 = 1$, and so the ringlet masses m_r and radii r quoted below are in units of M and r_0 . Converting code output from natural units to physical units requires choosing physical values for M and r_0 and multiplying accordingly, and when this text does so it assumes Saturn’s mass $M = 5.68 \times 10^{29}$ gm and a characteristic ring radius $r_0 = 1.0 \times 10^{10}$ cm. Simulation time t is in units of $T_{\text{orb}}/2\pi$ where $T_{\text{orb}} = 2\pi\sqrt{r_0^3/GM}$ is the orbit period at r_0 , so divide simulation time t by 2π and then multiply by T_{orb} to convert simulation time from natural to physical units. The simulated particles’ motions during the drift step are also sensitive to the J_2 portion of the primary’s non-spherical gravity component (see Appendix B), and all simulations adopt Saturn-like values of $J_2 = 0.01$ and $R_p = r_0/2$ where R_p is the planet’s mean radius.

Initially all particles are assigned to various streamlines across the simulated ringlet. A streamline is a closed eccentric path around the primary, and the N_p particles in a given streamline are initially assigned a common semimajor axis a and eccentricity e and are distributed uniformly in longitude. Most of the simulations described below employ only $N_s = 2$ streamlines, so that the model output can be benchmarked against theoretical treatments that also treat the ringlet as two gravitating rings (e.g. [Borderies et al. 1983](#)). But the following also performs a few higher-resolution simulations using $N_s = 11$ streamlines, to demonstrate that the $N_s = 2$ treatment is perfectly adequate and reproduces all the relevant dynamics. All simulations use $N_p = 241$ particles per streamline, and the total number of particles is $N_s N_p$. Note that the assignment of particles to a given streamline is merely for labeling purposes, as particles are still free to wander in response to the ring’s internal forces, namely, ring gravity and viscosity. But as [Hahn & Spitale \(2013\)](#) as well as this work shows, the simulated ring stays coherent and highly organized throughout the simulation such that particles on the same streamline do not pass each other longitudinally, nor do they cross adjacent streamlines. Because the simulated ringlet stays highly organized, there is no radial or longitudinal mixing of the ring particles, and simulated particles preserve their streamline membership over time.

4. N-BODY SIMULATIONS OF VISCOUS GRAVITATING RINGLETS

This Section describes a suite of N-body simulations of narrow viscous gravitating planetary ringlets, to highlight the range of initial ringlet conditions the do evolve into a self-confining state, and those that do not.

4.1. *nominal model*

Figure 1 shows the semimajor axis evolution of what is referred to as the nominal model since this ringlet readily evolves into a self-confining state. The simulated ringlet is composed of $N_s = 2$ streamlines having $N_p = 241$ particles per streamline, and the integrator timestep is $\Delta t = 0.5$ in natural units, so the integrator samples the particles’ orbits $2\pi/\Delta t \simeq 13$ times per orbit, and this ringlet is evolved for 4.7×10^3 orbits, which requires 15 minutes execution time on a 5 year old laptop. The ringlet’s mass is $m_r = 5 \times 10^{-10}$, its shear viscosity is $\nu_s = 2.5 \times 10^{-12}$, and its bulk viscosity is $\nu_b = \nu_s$. The ringlet’s initial radial width is $\Delta a_0 = 3 \times 10^{-4}$, its initial eccentricity is $e = 0.01$, and its eccentricity gradient is initially zero. A convenient measure of time is the ringlet’s viscous radial spreading timescale

$$\tau_\nu = \frac{\Delta a_0^2}{12\nu_s}, \quad (1)$$

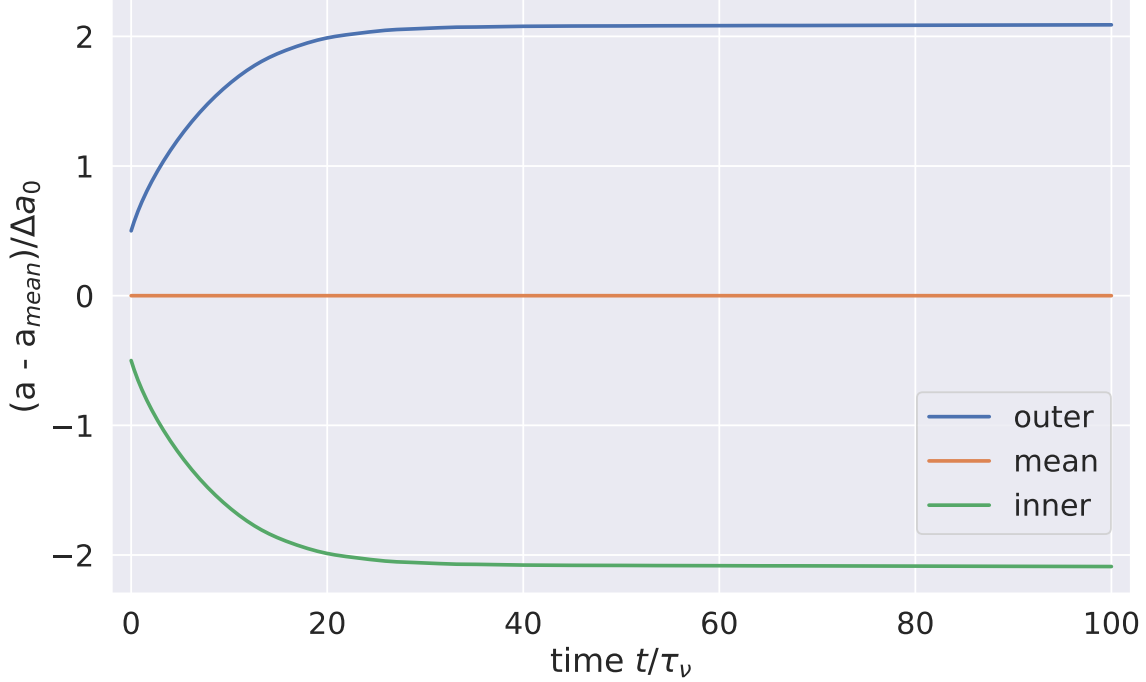


Figure 1. Evolution of the nominal ringlet’s semimajor axes a versus time t in units of the ringlet’s viscous time τ_ν . This ringlet is composed of $N_s = 2$ streamlines, and the outer (blue) and inner (green) streamlines’ semimajor axes are plotted relative to their mean a_{mean} , and displayed in units of the ringlet’s initial width $\Delta a_0 = 3 \times 10^{-4}$ in natural units (*i.e.* $G = M = r_0 = 1$). The simulated ringlet has total mass $m_r = 5 \times 10^{-10}$, shear viscosity $\nu_s = 2.5 \times 10^{-12}$, and initial eccentricity $e = 0.01$, and see Section 4.1 to convert m_r , a and ν_s from natural units to physical units.

which can be inferred from Eqn. (2.13) of [Pringle \(1981\)](#). This simulation’s viscous timescale is $\tau_\nu = 3.0 \times 10^3$ in natural units or $\tau_\nu/2\pi = 4.8 \times 10^2$ orbital periods. If this ringlet were orbiting Saturn at $r_0 = 1.0 \times 10^{10}$ cm then the simulated ringlet’s physical mass would be $m_r = 2.8 \times 10^{20}$ gm which is equivalent to the mass of a 41 km radius iceball assuming a volume density $\rho = 1$ gm/cm³, and the ringlet’s initial radial width would be $\Delta a_0 = 3 \times 10^{-4} r_0 = 30$ km. This ringlet’s orbit period would be $T_{\text{orb}} = 2\pi\sqrt{r_0^3/GM} = 9.0$ hours in physical units, so the ringlet’s viscous timescale is $\tau_\nu = 12$ years, and so its shear viscosity is $\nu_s = \Delta a_0^2/12\tau_\nu = 4.8 \times 10^4$ cm²/sec when evaluated in physical units. This ringlet’s initial surface density would be $\sigma = m_r/2\pi r_0 \Delta a_0 = 1500$ gm/cm², but Figs. 1–2 show that shrinks by a factor of 4 as the ringlet’s sememajor axis width Δa grows via viscous spreading until it settles into the self-confining state at time $t \sim 20\tau_\nu$. This so-called nominal ringlet is probably overdense and overly viscous compared to known planetary ringlets, but that is by design so that the simulated ringlet quickly settles into the self-confining state. Section XX also shows how outcomes scale when a wide variety of alternate initial masses, orbits, and viscosities are also considered.

Figure 3 shows that the outer streamline’s eccentricity initially grows at the expense of the inner streamline’s, and this is a consequence the self-gravitating ringlet’s secular perturbations of itself, which is also demonstrated in Appendix C. Figure 4 shows the ringlet’s eccentricity difference

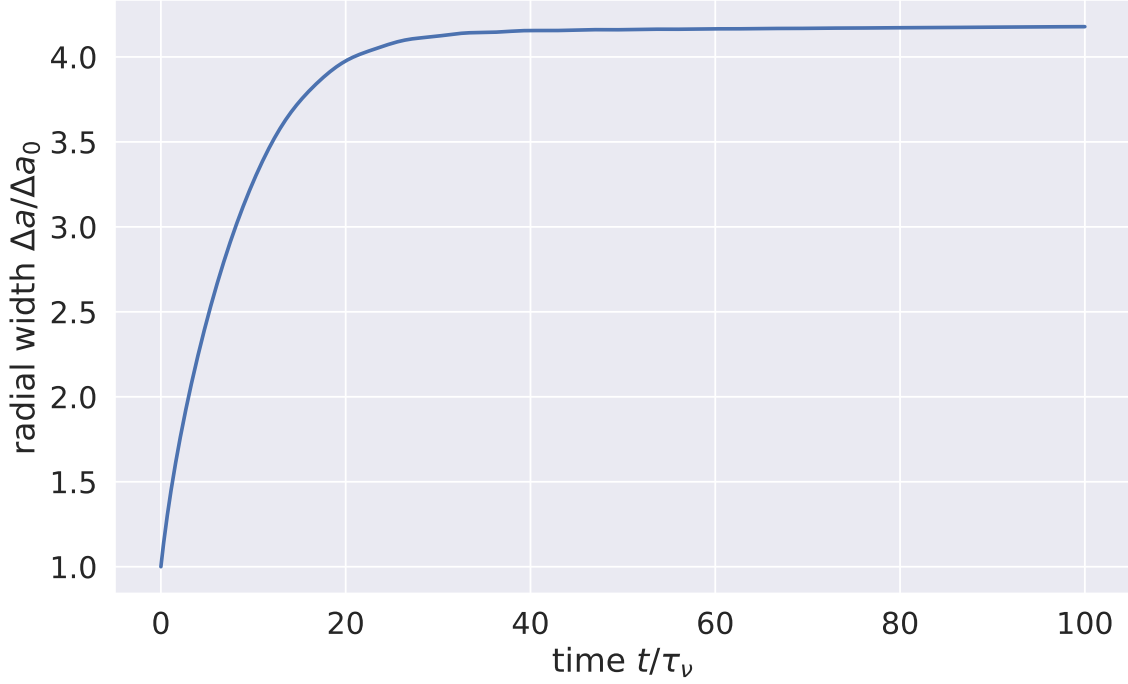


Figure 2. The nominal ringlet's semimajor axis width $\Delta a = a_{\text{outer}} - a_{\text{inner}}$ over time and in units of its initial radial width Δa_0 .

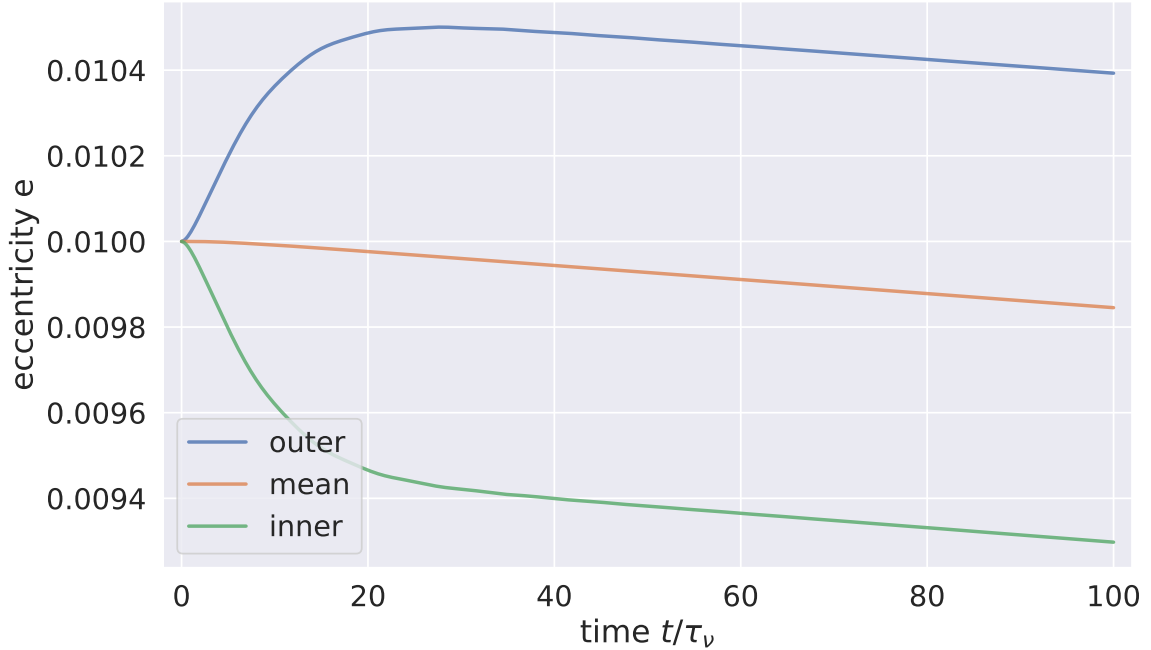


Figure 3. The nominal ringlet's eccentricity evolution.

$\Delta e = e_{\text{outer}} - e_{\text{inner}}$ and longitude of periapse difference $\Delta \tilde{\omega} = \tilde{\omega}_{\text{outer}} - \tilde{\omega}_{\text{inner}}$, which both settle into equilibrium values after the ringlet arrives in the self-confining state.

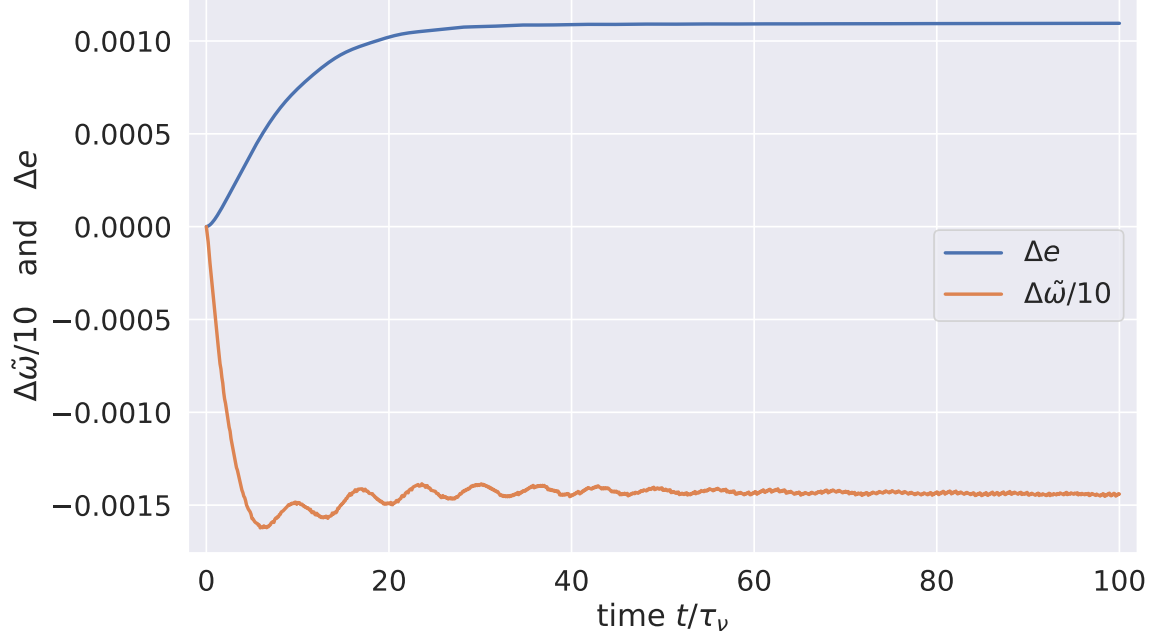


Figure 4. The nominal ringlet's eccentricity difference $\Delta e = e_{\text{outer}} - e_{\text{inner}}$ and longitude of periapse difference $\Delta\tilde{\omega} = \tilde{\omega}_{\text{outer}} - \tilde{\omega}_{\text{inner}}$ in radians and divided by 10.

Figure 5 shows the radii of the ringlet's two streamlines plotted versus their relative longitude $\varphi = \theta - \tilde{\omega}_{\text{inner}}$ at time $t = 100\tau_v$ when the simulation ends. In all simulations examined here, the ringlet's periapse twist $\Delta\tilde{\omega} = \tilde{\omega}_{\text{outer}} - \tilde{\omega}_{\text{inner}}$ is negative, so the outer streamline's longitude of periapse $\tilde{\omega}$ trails the inner streamline's, which in turn causes the streamlines' separations along the ringlet's pre-periapse side (where $\varphi < 0$) to be smaller than at post-periapse ($\varphi > 0$). Which makes the ringlet's surface density asymmetric, with maximum surface density occurring just prior to periapse, see Figs. 5–7.

4.2. angular momentum and energy fluxes

It is convenient to recast these orbit element differences as dimensionless gradients

$$e' = a \frac{de}{da} \quad \text{and} \quad \tilde{\omega}' = ea \frac{d\tilde{\omega}}{da} \quad (2)$$

as these are the terms that contribute to the nonlinearity parameter of [Borderies et al. \(1983\)](#):

$$q = \sqrt{e'^2 + \tilde{\omega}'^2}. \quad (3)$$

See also Fig. 8 which plots the nominal ringlet's dimensionless eccentricity gradient e' , dimensionless periapse twist $\tilde{\omega}'$, and nonlinearity parameter q versus time. All simulations examined here have $|\tilde{\omega}'| \ll |e'|$ so that $q \simeq |e'|$, and all simulated self-confining ringlets have a positive eccentricity gradient and a negative periapse twist such that the outer ringlet's periapse trails the inner ringlet's, consistent with the findings of [Borderies et al. \(1983\)](#).

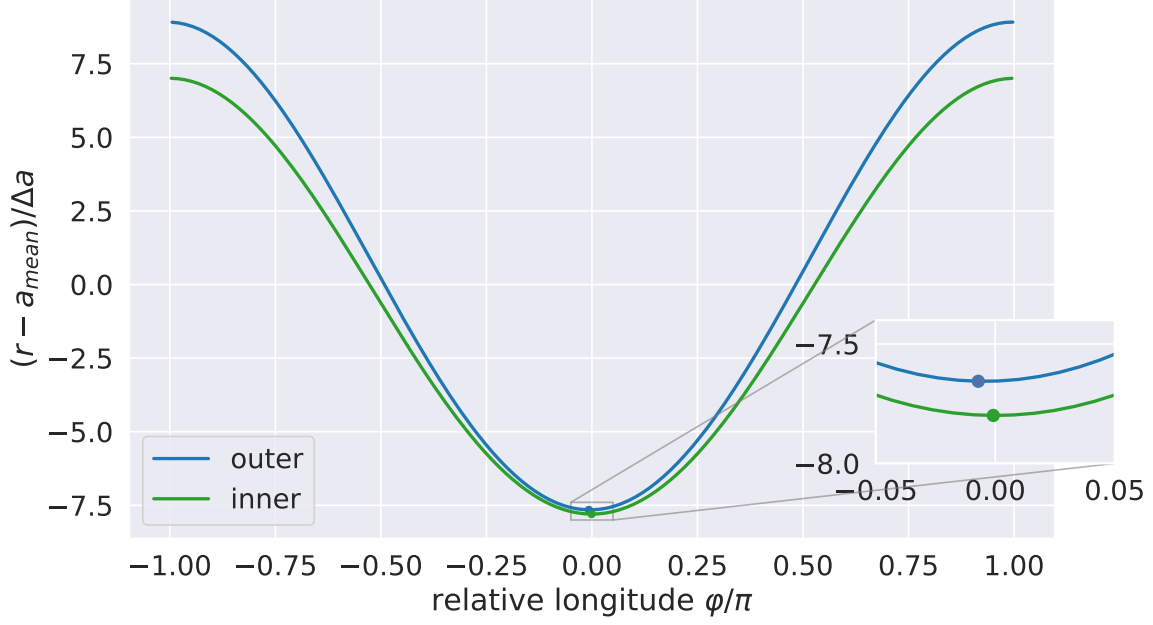


Figure 5. The radii of the nominal ringlet's streamlines are plotted versus relative longitude $\varphi = \theta - \tilde{\omega}$ at time $t = 100\tau_\nu$, with inset plot showing outer streamline's longitude of periapse $\tilde{\omega}$ trailing the inner streamline's.

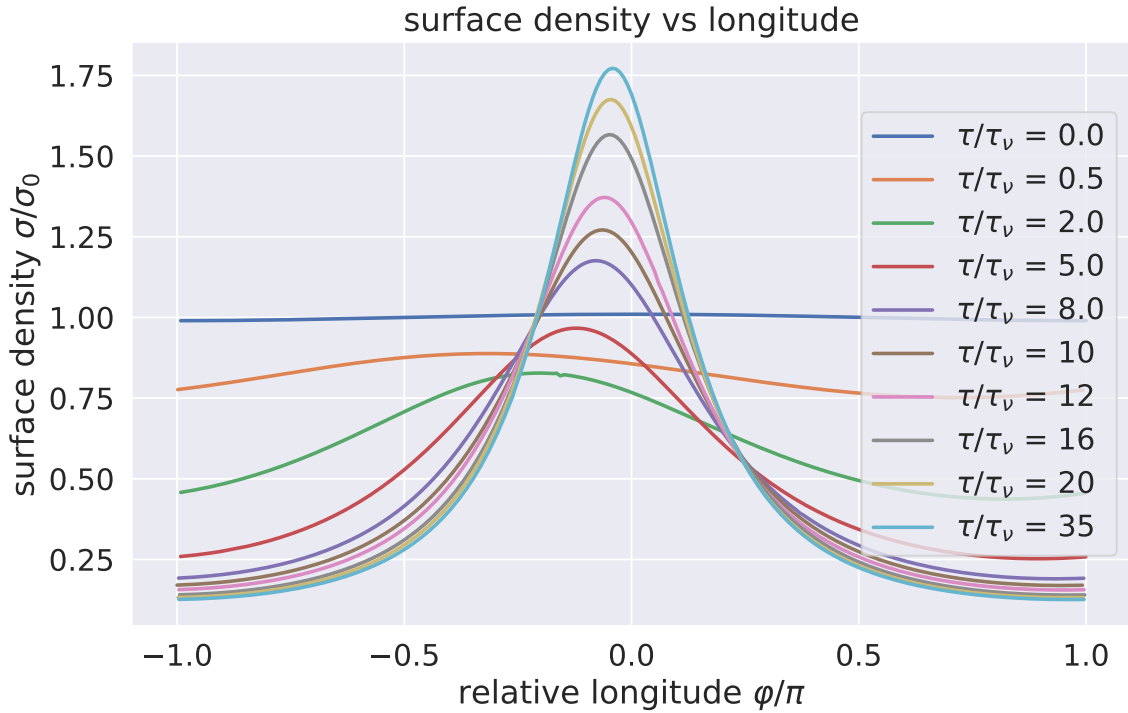


Figure 6. Nominal ringlet's surface density $\sigma(\varphi)$ is plotted versus relative longitude φ at selected times t . That the ringlet's surface density maxima occurs just before peripase is due to the ringlet's negative periapse twist $\Delta\tilde{\omega} = \tilde{\omega}_{\text{outer}} - \tilde{\omega}_{\text{inner}} < 0$.

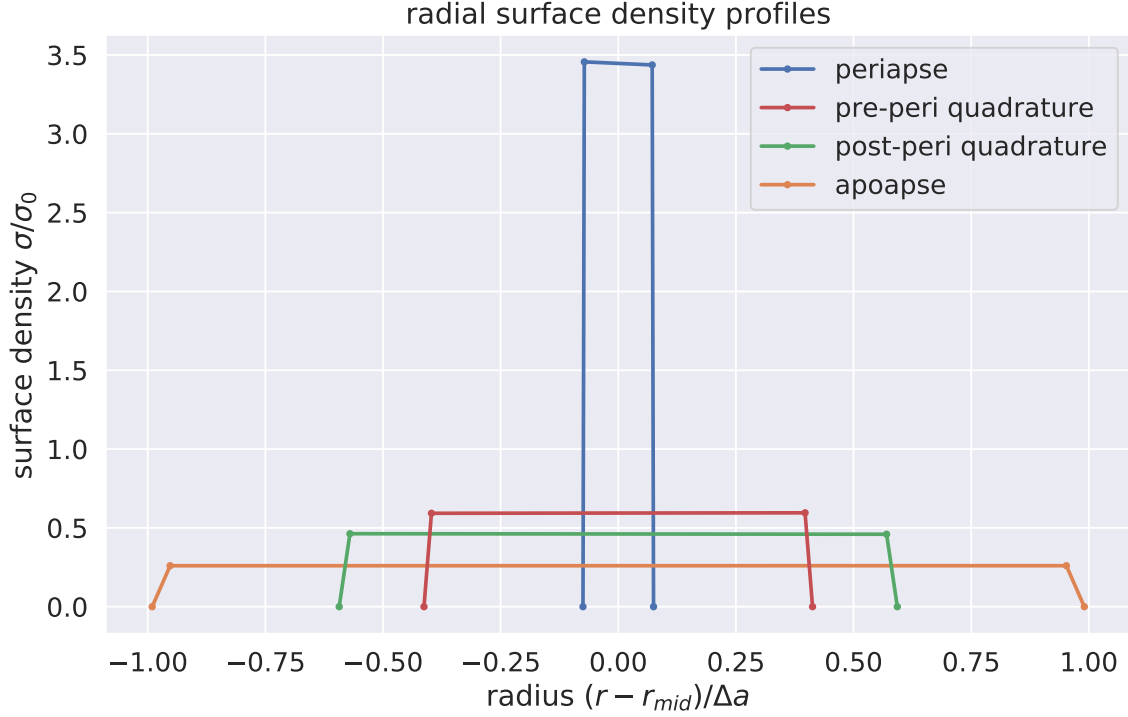


Figure 7. Radial profiles of the nominal ringlet’s surface density $\sigma(\varphi)$ at time $t/\tau_\nu = 100$ when the ringlet is self-confining. Each surface density profile is plotted versus radial distance r relative to r_{mid} , which is the ringlet’s midpoint along relative longitude $\varphi = \theta - \tilde{\omega}$, with those radial distances $r - r_{mid}$ measured in units of the ringlet’s final semimajor axis width Δa , and surface density is shown in units of the ringlet’s longitudinally-averaged surface density σ_0 . Radial surface density profiles are plotted along the ringlet’s periapse ($\varphi = 0$, blue curve), which is where the ringlet’s streamlines are most concentrated and surface density σ is greatest due to the ringlet’s eccentricity gradient e' , at quadrature ($|\varphi| = \pi/2$, green curve), and at apoapse ($|\varphi| = \pi$, orange curve) where streamlines have their greatest separation and ringlet surface density is lowest. This ringlet’s surface density contrast, between periapse and apoapse, is 14.

4.3. viscous angular momentum transport

The ringlet’s evolution is readily understood when the ringlet’s viscous flux of angular momentum is considered; that flux is

$$F_\nu(r, \theta) = -\nu_s \sigma r^2 \frac{\partial \omega}{\partial r} \quad (4)$$

(Hahn & Spitale 2013) when written as a function of spatial coordinates and angular velocity $\omega = \dot{\theta}$ (Eqn. XX). If we consider a small arc of ring material of transverse length $d\ell$, then $F_\nu d\ell$ would be the torque that arc exerts on ring matter just exterior, due to viscous friction, so that is the rate that friction transmits angular momentum radially across that arc. And when F_ν is evaluated along a single eccentric streamline of semimajor axis a , the above simplifies to

$$F_\nu(a, \varphi) = F_{\nu,0} \frac{1 - \frac{4}{3}e' \cos \varphi}{(1 - e' \cos \varphi)^2} \quad (5)$$

where $\varphi = \theta - \tilde{\omega}$ is the longitude relative to periapse (see Borderies et al. 1982 and Appendix D) and $F_{\nu,0} = \frac{3}{2}\nu_s \sigma_0 a \Omega$ is the viscous angular momentum flux through a circular streamline of semimajor

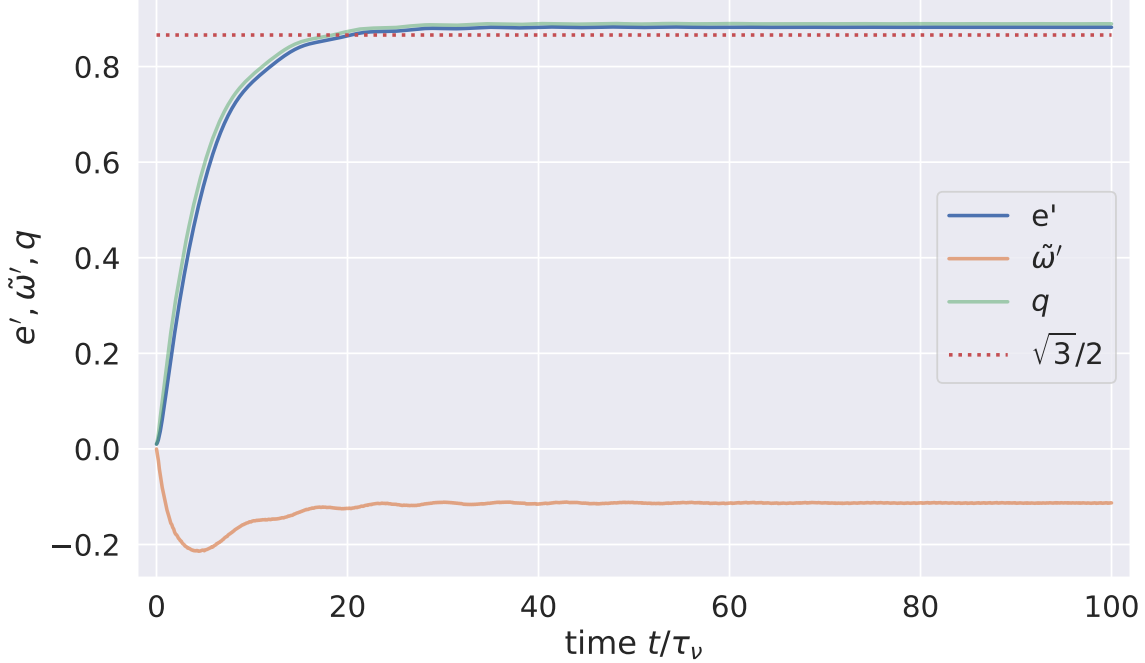


Figure 8. The nominal ringlet’s dimensionless eccentricity gradient $e' = a\Delta e/\Delta a$ (blue curve), dimensionless periapse twist $\tilde{\omega}' = ea\Delta\tilde{\omega}/\Delta a$ (orange), and nonlinearity parameter $q = \sqrt{e'^2 + \tilde{\omega}'^2}$ (green) versus time t/τ_ν . Dotted red line is the threshold for self-confinement in a non-gravitating ringlet, $e' = \sqrt{3}/2 \simeq 0.866$

axis a and angular speed $\Omega(a)$ and assuming $|\tilde{\omega}'| \ll e'$ so that $q \simeq e'$. Integrating the above around the streamline’s circumference then yields its angular momentum luminosity,

$$L_\nu(a) = \oint F_\nu(e', \varphi) r d\varphi = L_{\nu,0} \frac{1 - \frac{4}{3}e'^2}{(1 - e'^2)^{3/2}}, \quad (6)$$

which is the torque that one streamline exerts on its exterior neighbor due to viscous friction (Borderies et al. 1982 and Appendix D), where $L_{\nu,0} = 3\pi\nu_s\sigma_0a^2\Omega$ viscous angular momentum luminosity of a circular streamline.

Borderies et al. (1982) examine angular momentum transport through a viscous eccentric but non-gravitating ringlet, and use Eqns. (5–6) to show that this transport has three regimes distinguished by the ringlet’s e' :

1. $e' < 3/4$, so the ringlet’s viscous angular momentum flux $F_\nu(\varphi) > 0$ at all ringlet longitudes θ . The ringlet’s viscous angular momentum luminosity $L_\nu > 0$, so viscous friction transports angular momentum radially outwards, so inner ring matter evolves to smaller orbits while exterior ring matter evolves outwards, and the ringlet spreads radially.
2. $3/4 \leq e' < \sqrt{3}/2$. In this regime there is a range of longitudes θ where the viscous angular momentum flux is reversed such that $F_\nu(\varphi) < 0$. Nonetheless L_ν , which is the orbit-average of $F_\nu(\varphi)$, is positive and the ringlet still spreads radially, albeit slower than when $e' < 0.75$.
3. $e' \geq \sqrt{3}/2$. Viscous angular momentum flux reversal is complete such that $L_\nu < 0$, viscous friction transports angular momentum radially inwards, and the ringlet shrinks radially. But

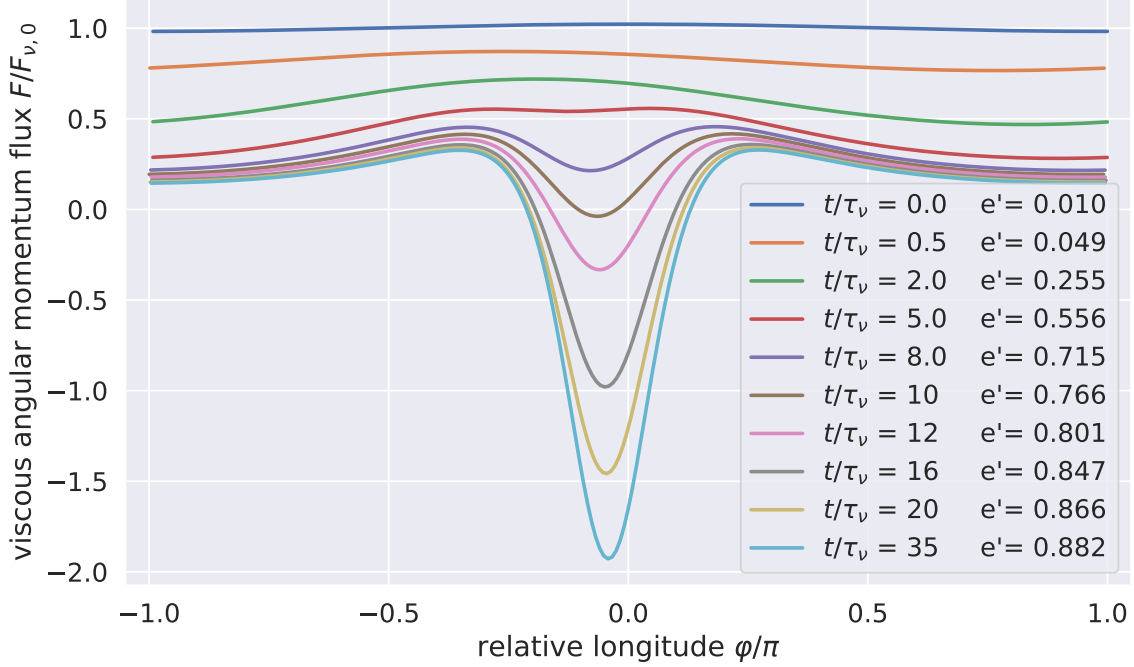


Figure 9. The nominal ringlet’s viscous angular momentum flux $F_\nu(\varphi)$, Eqn. (5), is plotted versus ringlet relative longitude $\varphi = \theta - \tilde{\omega}$ about the ringlet’s inner streamline at selected times t/τ_ν , with the ringlet’s eccentricity gradient e' also indicated.

if $e' = \sqrt{3}/2 \simeq 0.866$ then $L_\nu = 0$ and the ringlet’s radial evolution ceases, and the viscous but non-gravitating ringlet is self confining.

Note though that the nominal ringlet’s eccentricity gradient exceeds the $e' = \sqrt{3}/4 \simeq 0.866$ threshold (which is the dotted red line in Fig. 8) when it settles into self-confinement. This is due to the ringlet’s self-gravity, which also transports angular momentum radially through the ringlet.

Figure 9 shows the nominal ringlet’s viscous angular momentum flux F_ν versus relative longitude $\varphi = \theta - \tilde{\omega}$ at selected times t . Early in the ringlet’s evolution when time $t \leq 8\tau_\nu$ (blue, orange, green, red, and purple curves), the ringlet is in regime 1 since $e' < 0.75$ and $F_\nu(\varphi) > 0$ at all longitudes. But by time $t = 10\tau_\nu$ (brown curve), this ringlet’s eccentricity gradient exceeds 0.75, and angular momentum flux reversal $F_\nu(\varphi) < 0$ occurs near periaapse where $|\varphi| \simeq 0$ where the ringlet is most overdense due to its eccentricity gradient, see also Fig. 7; this ringlet is in regime 2 and its radial spreading is reduced by angular momentum flux reversal. And by time $t = 20\tau_\nu$ (yellow curve), this ringlet is seemingly in regime 3 since $e' = 0.866$, so one might expect the ringlet’s spreading to have stalled by now, but keep in mind that the above analysis ignores any transport of angular momentum via ringlet self-gravity. Figure 2 shows that this gravitating ringlet’s spreading has ceased soon after time $t \simeq 35\tau_\nu$, at which point $e' = 0.88$ (cyan curve), angular momentum flux reversal is nearly complete, with the ringlet’s total angular momentum luminosity $L = L_\nu + L_g = 0$ is very close to zero.

Figure 10 and Fig. 11 show that, when the ringlet is self-confining, it’s positive viscous angular momentum luminosity L_ν is nearly counterbalanced by it’s negative gravitational angular momentum luminosity L_g .

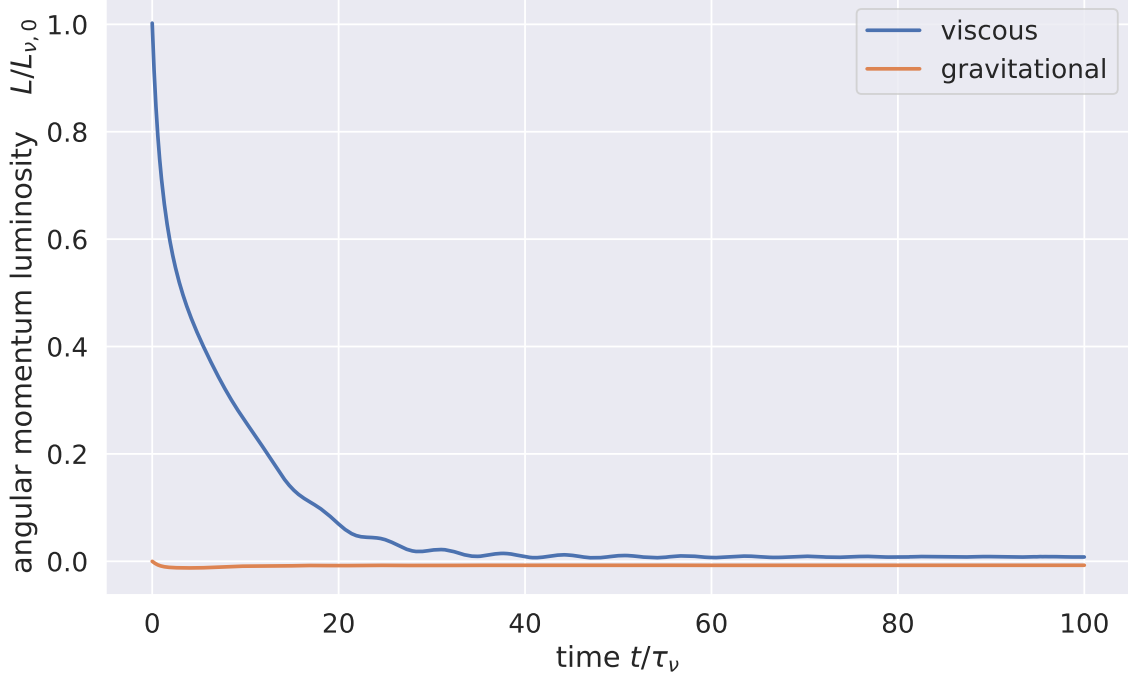


Figure 10. Nominal ringlet’s viscous angular momentum luminosity $L_{\nu}/L_{\nu,0}$ (blue curve) versus time t/τ_{ν} , where $L_{\nu,0}$ is circular ring’s viscous angular momentum luminosity, as well as the ringlet gravitational angular momentum luminosity L_g (orange curve) in units of $L_{\nu,0}$.

The ringlet’s viscous F_{ν} and gravitational F_g angular momentum fluxes are shown Fig. 12, and that figure shows how viscosity and self-gravity both transport angular momentum radially through the self-confining ringlet at various longitudes φ . To compute the ringlet’s gravitational angular momentum flux F_g , note that the gravitational torque that is exerted on a small streamline segment of mass δm due to the ring material orbiting interior is $T_g = \delta m r A_{g,\theta}^1 = F_g \delta \ell$, where A_g^1 is the so-called one-sided gravitational acceleration *i.e.* the gravitational acceleration that segment δm experiences due to the ring matter orbiting interior to it. The tangential component of that one-sided acceleration is $A_{g,\theta}^1$, so for a segment whose length is $\delta \ell$, its gravitational angular momentum flux is $F_g = \lambda r A_{g,\theta}^1$ where $\lambda = m_r / N s / 2\pi a$ is the streamline’s linear density. The streamline’s total gravitational angular momentum luminosity is

$$L_g(a) = \oint F_g(\varphi) r d\varphi = \oint \lambda r^2 A_g^1(\varphi) d\varphi, \quad (7)$$

which is the orange curve in Figs. 10–11 and is the gravitational torque that the inner streamline exerts on the outer. *To do: explain why L isn’t precisely zero when ringlet is seemingly self-confining...*

That figure shows that ringlet viscosity transports angular momentum inwards *i.e.* $F_{\nu}(\varphi) < 0$ near periapse, and outwards with $F_{\nu}(\varphi) > 0$ at all other longitudes. Which is rather distinct from the ringlet’s gravitational transport, which has $F_g(\varphi) < 0$ as ring-matter travels towards periapse, and outwards $F_g(\varphi) > 0$ after periapse. Despite these spatial differences, the influence of both forces still sum to zero in the orbit-integrated sense *i.e.* $\oint (F_{\nu} + F_g) r d\varphi = 0$ after the ringlet has settled into the self-confining state.

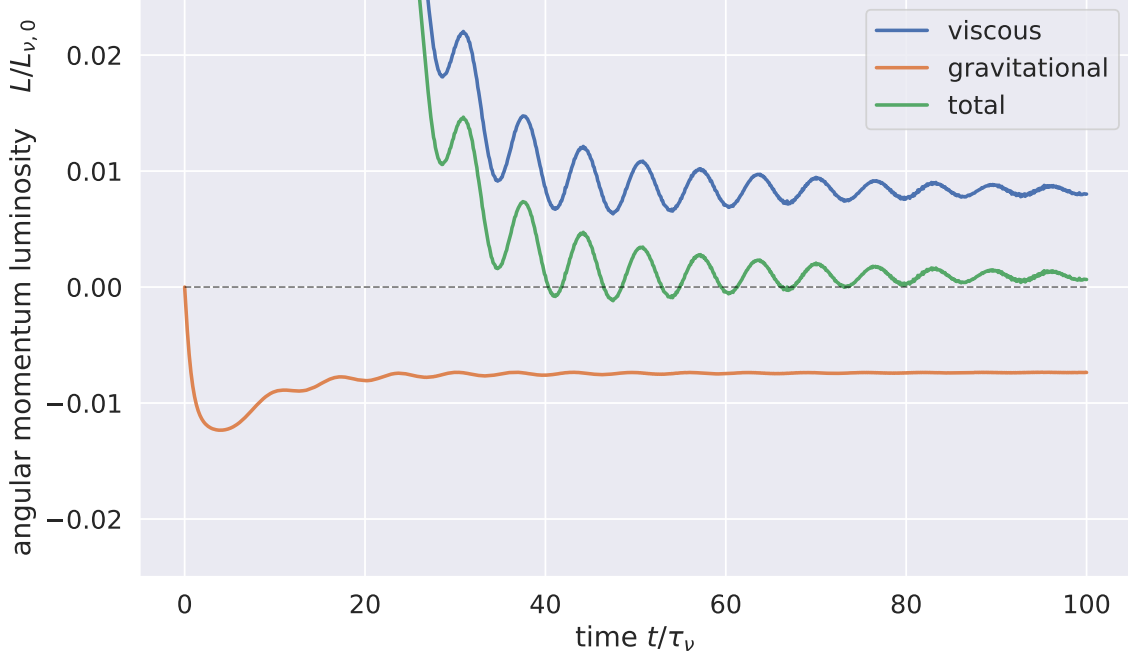


Figure 11. Figure 10 is replotted to highlight that the ringlet’s viscous angular momentum luminosity L_ν (blue curve) always stays positive (indicating that the viscous transport of angular momentum is radially outwards) which is nearly balanced by the ringlet’s negative (*i.e.* inwards) gravitational angular momentum luminosity L_g (orange) after time $t \gg 35\tau_\nu$.

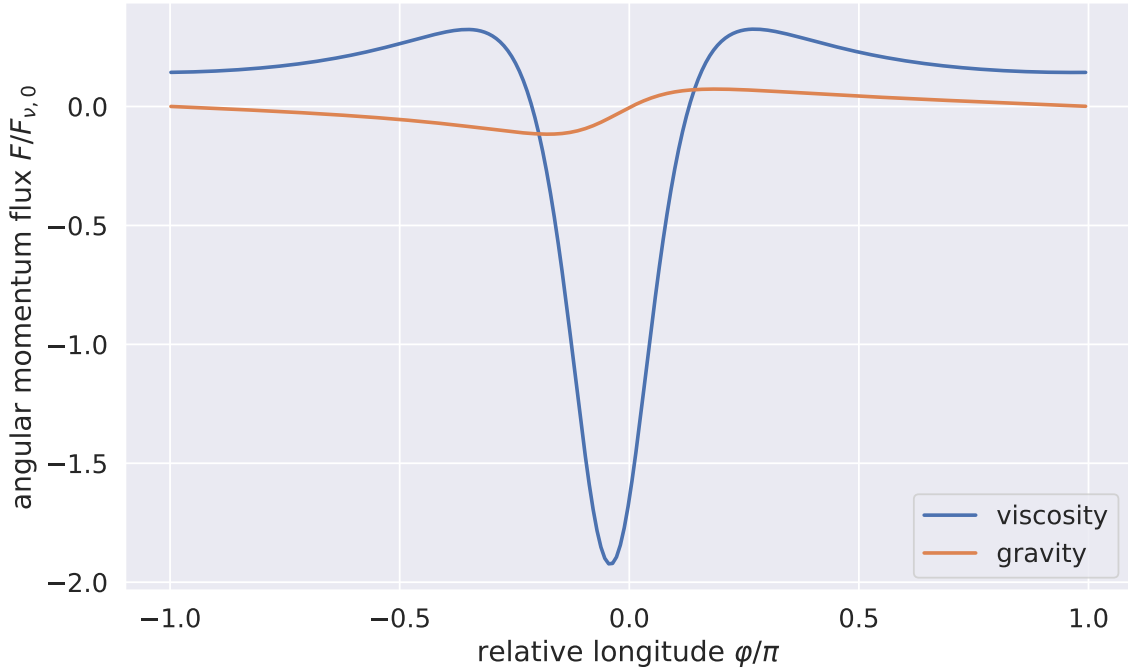


Figure 12. The nominal ringlet’s viscous angular momentum flux $F_\nu(\phi)$ (blue curve), plotted in units of $F_{\nu,0}$ and versus relative longitude ϕ as the simulation’s end at time $t = 100\tau_\nu$, as well as the ringlet’s gravitational angular momentum flux $F_g(\phi)$ in units of $F_{\nu,0}$ (orange curve).

Ring viscosity and self-gravity can also transport energy across the ring, and that is assessed in Appendix E.

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APPENDIX

A. APPENDIX A

Derive the more accurate drift step used by `epi_int_lite`...

B. APPENDIX B

Compare `epi_int_lite` to theoretical predictions

C. APPENDIX D

This examines the viscous evolution of a narrow eccentric non-gravitating ringlet that is identical to the nominal ringlet of Section 4.1 but with ringlet self-gravity neglected and $J_2 = 0$. As the orange curve in Fig. 13 shows, the non-gravitating ringlet's radial width Δa grows steadily over time due to ringlet viscosity, long after the nominal self-gravitating ringlet (blue curve) has settled into the self-confining state by time $t \sim 15\tau_\nu$. This is due to the ringlet's secular gravitational perturbations of itself, which tends to excites the ringlet's outer streamline's eccentricity at the expense of the inner streamline (see Fig. 3) until the ringlet eccentricity gradient e' (blue curve in Fig. 14) grows beyond the limit required for complete angular momentum flux reversal that results in the ringlet's radial confinement (dotted line). Note that viscosity also excites the non-gravitating ringlet's eccentricity gradient some (orange curve), but not sufficiently to halt the ringlet's viscous spreading.

D. APPENDIX E

This Appendix will use the orbit elements derived in Appendix A to derive Eqn. 5 from 4, and then Eqn. (6).

E. APPENDIX F

Viscous and gravitational energy transport...

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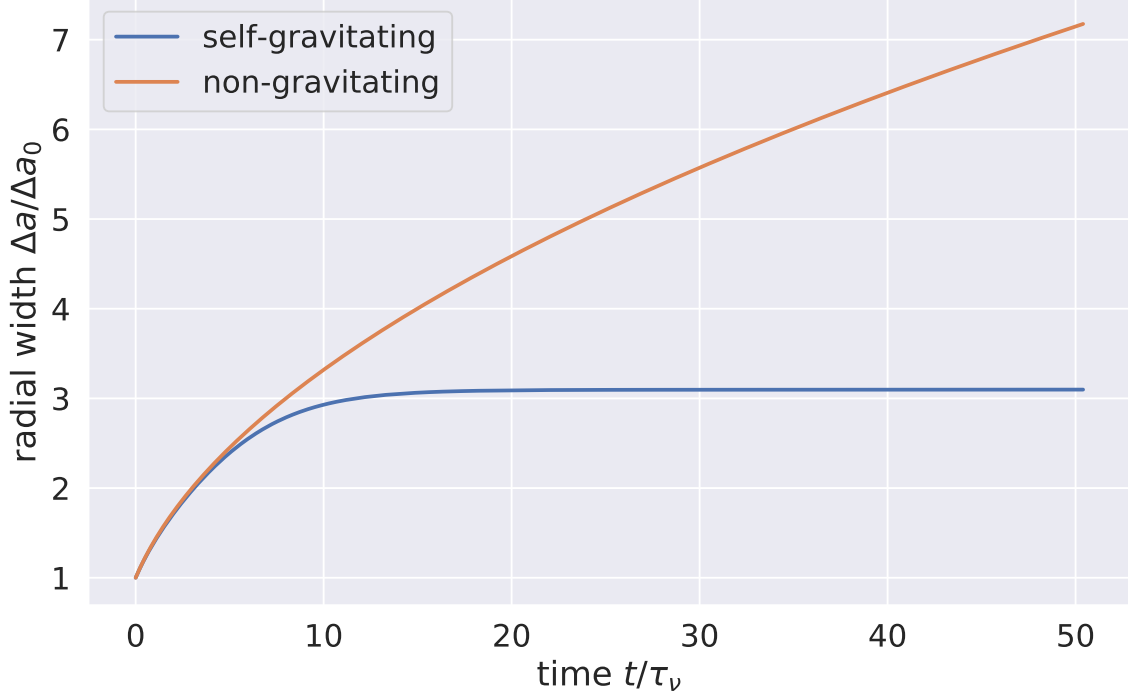


Figure 13. Blue curve is the nominal ringlet’s semimajor axis width Δa versus time t , and this ringlet’s radial spreading ceases by time $t \sim 15\tau_v$ when it’s self-gravity has excited the ringlet’s eccentricity gradient e' sufficiently; see blue curve in Fig. 14. Orange curve shows that the non-gravitating ringlet’s Δa grows without limit due to the ringlet’s much lower eccentricity gradient. Note that planetary oblateness would cause the non-gravitating streamlines to precess differentially and eventually cross when $J_s > 0$, so the non-gravitating simulation also sets $J_2 = 0$ to avoid differential precession.

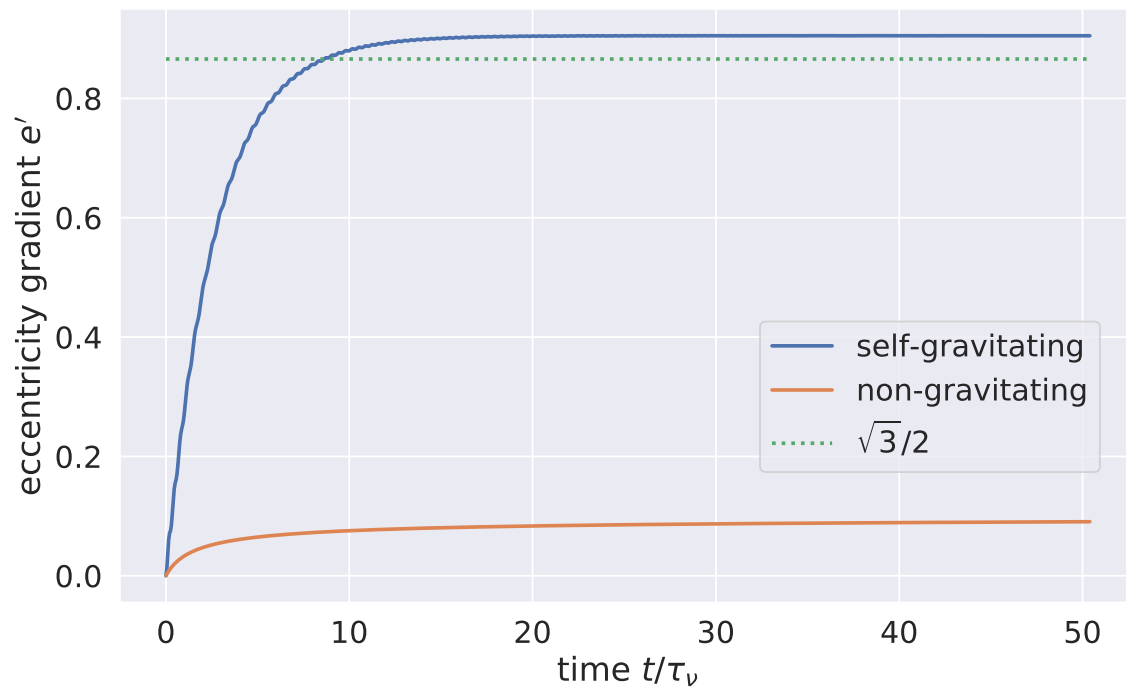


Figure 14. blah