Todo list

OBSERVATIONS OF PWNE WITH THE FERMI GAMMA-RAY SPACE TELESCOPE

A DISSERTATION
SUBMITTED TO THE DEPARTMENT OF PHYSICS
AND THE COMMITTEE ON GRADUATE STUDIES
OF STANFORD UNIVERSITY
IN PARTIAL FULFILLMENT OF THE REQUIREMENTS
FOR THE DEGREE OF
DOCTOR OF PHILOSOPHY

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April 2013

I certify that I have read this dissertation and that, in my opinion, it is fully adequate in scope and quality as a dissertation for the degree of Doctor of Philosophy.

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(Roger Romani)

Approved for the University Committee on Graduate Studies

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List of Acronyms

1FHL the first *Fermi* hard-source list. 126

2CG the second COS-B catalog. 9

2FGL the second *Fermi* catalog. 22, 38, 122, 127

2PC the second *Fermi* pulsar catalog. 23, 124, 125

3EG the Third EGRET Catalog. 11

ACD Anti-Coincidence Detector. vii, 12

arcsec second of arc. 32

BPL broken-power law. 40

CGRO the Compton Gamma Ray Observatory. 10

CGS the Centimetre-Gram-Second System of Units. 40, 41

ECPL exponentially-cutoff power law. 40, 41

EGRET the Energetic Gamma Ray Experiment Telescope. xii, 9–11

ESA the European Space Agency. 9

FWHM full width at half maximum. 8

GBM Gamma-ray Burst Monitor. vii, 12

IACT Imaging air Cherenkov detector. 48, 119, 122, 126

IC inverse Compoton. 6, 19, 20, 31, 32, 35, 126

LAT the Large Area Telescope. iv, v, vii, viii, 12, 22, 23, 122, 123, 125–127, 129, 131

MIT the Massachusetts Institute of Technology. 6, 13

MSC massive star cluster. 125

MSP millisecond pulsar. 27

NASA the National Aeronautics and Space Administration. 9, 10

NRL the Naval Research Laboratory. 13

NS neutron star. 13, 25, 26, 30

OSO-3 the Third Orbiting Solar Observatory. xii, 8, 9, 22

PL power law. 40, 41

PSF point spread function. 126

PWN pulsar wind nebula. iv, v, xiii, 1, 6, 14, 25, 29, 31–35, 122, 123, 125–129, 131

SA solid angle. 42, 43

SAS-2 the second Small Astronomy Satellite. 9, 10

SNR supernova remnant. 29, 32, 41

UNID unidentified source. 125, 126

VHE very high energy. ix, xi, 122–129, 133

Chapter 9

Population Study of LATs-detected PWNe

This chapter is based the second part of the the paper "Constraints on the Galactic Population of TeV Pulsar Wind Nebulae using Fermi Large Area Telescope Observations" by Acero et al which is currently in prep.

In Chapter 6, we search for new spatially-extended *Fermi* sources and found that spatial extenion was an important characteristic for detecting new pulsar wind nebulae (PWNe). In the process, we discovered three new γ-ray emitting PWNs. In Chapter 7, we then searched in the off-peak phase interval of the Large Area Telescope (LAT)-detected pulsars for new pulsar wind nebula and discovered 3C 58. Finally, in Chapter 8 we searched in the regions surrounding PWNs candidates detected at TeV energies for GeV-emitting PWNs 4 new PWNe candidates (HESS J1119−614, HESS J1303−631, HESS J1420−607, and HESS J1841−055) and 1 new PWN (HESS J1356−645)

In this chapter, we take the population of γ-ray emitting PWNs and PWNs candidates

9.1 Summary of the PWNe detected by the LATs

9.2 The Evolution of γ-ray Emitting PWNe with the Properties of their Pulsars

Table 9.1.

Source	$\dot{E}$ ($\mathrm{erg\,s^{-1}}$)	τ (kyr)	Distance (kpc)	Flux TeV	Flux Xray
VERJ0006+727	4.5e+35	13.9	1.4 +0.3 -0.3	···	···
Crab	4.6e+38	1.2	2.0 +0.5 -0.5	80.0 +/- 16.49242	21000.0 +/- 4200.0
MGROJ0631+105	1.7e+35	43.6	1.0 +0.2 -0.2	···	···
MGROJ0632+17	3.2e+34	342	0.2 +0.01 -0.02	···	···
Vela-X	6.9e+36	11.3	0.3 +0.02 -0.02	79.0 +/- 21.2132	53.9 +/- 17.8
HESSJ1018-589	2.6e+36	21	2.9 +1.6 -0.06	0.9 +/- 0.4	···
HESSJ1023-577	1.1e+37	4.6	2.8 +0.0 -0.0	4.82411 +/- 1.70053	···
HESSJ1026-582	8.4e+35	90	2.3 +0.3 -0.3	3.73537 +/- 2.18879	···
HESSJ1119-614	2.3e+36	1.6	8.4 +0.4 -0.4	2.3188 +/- 1.15666	···
HESSJ1303-631	1.7e+36	11	6.65 +1.2 -1.1	26.7 +/- 1.0	0.16 +/- 0.03
HESSJ1356-645	3.1e+36	7.3	2.5 +0.53 -0.44	6.66231 +/- 3.71138	0.0604 +/- 0.012
HESSJ1418-609	4.9e+36	1.03	1.6 +0.7 -0.7	3.41 +/- 1.82	3.11 +/- 0.147
HESSJ1420-607	1.0e+37	13	5.6 +0.91 -0.85	14.5 +/- 3.3	1.3 +/- 0.26
HESSJ1458-608	9.1e+35	64.7	4.0 +0.0 -0.0	2.45224 +/- 0.68099	···
HESSJ1514-591	1.7e+37	1.56	4.2 +0.6 -0.6	20.3 +/- 4.25	28.6 +/- 5.72
HESSJ1554-550	···	18	7.5 +1.3 -1.3	···	3.07 +/- 1.0
HESSJ1616-508	1.6e+37	8.13	6.8 +0.7 -0.7	21.0 +/- 5.0	4.2 +/- 0.84
HESSJ1632-478	3.0e+36	20	3.0 +0.0 -0.0	14.8003 +/- 4.71778	0.43 +/- 0.08
HESSJ1640-465	4.0e+36	···	8.6 +0.0 -0.0	5.46577 +/- 1.23707	0.46 +/- 0.092
HESSJ1646-458B	2.1e+35	110	5.7 +0.7 -0.7	3.0 +/- 0.8	···
HESSJ1702-420	3.4e+35	55	4.8 +0.6 -0.5	9.1 +/- 3.4	0.006 +/- 0.0
HESSJ1708-443	3.4e+36	17.5	2.3 +0.3 -0.3	22.88 +/- 6.96	0.72 +/- 0.0339
HESSJ1718-385	1.3e+36	89.5	4.24 +0.4 -0.4	4.3 +/- 1.6	0.14 +/- 0.028
HESSJ1804-216	2.2e+36	16	3.8 +0.5 -0.4	11.8 +/- 2.4	0.068 +/- 0.0136
HESSJ1809-193	1.8e+36	51.3	3.5 +0.4 -0.4	19.0 +/- 5.66	0.23 +/- 0.046
HESSJ1813-178	6.8e+37	5.4	···	4.97349 +/- 0.6491	···
HESSJ1818-154	2.3e+33	9	7.8 +1.4 -1.6	1.29742 +/- 0.90339	···
HESSJ1825-137	2.8e+36	21	4.12 +0.4 -0.4	61.0 +/- 13.89	0.44 +/- 0.088
HESSJ1831-098	1.1e+36	128	4.0 +0.4 -0.4	5.08 +/- 0.58	···
HESSJ1833-105	3.4e+37	4.85	4.7 +0.4 -0.5	2.4 +/- 1.21	40.0 +/- 0.0
HESSJ1837-069	5.5e+36	2.23	6.6 +0.9 -0.9	22.9 +/- 8.58	0.639 +/- 0.243
HESSJ1841-055	5.9e+36	4.97	1.3 +0.0 -0.0	23.47835 +/- 3.43665	···
HESSJ1846-029	···	···	···	8.97 +/- 1.51	29.4 +/- 1.39
HESSJ1848-018	···	···	6.0 +0.0 -0.0	4.32062 +/- 0.9902	···
HESSJ1849-000	9.8e+36	42.9	7.0 +0.0 -0.0	2.12 +/- 0.41	0.9 +/- 0.2
HESSJ1857+026	4.6e+36	20.6	9.0 +1.2 -1.2	18.41 +/- 2.88	···
MGROJ1908+06	2.8e+36	19.5	3.2 +0.3 -0.3	11.9 +/- 4.77	···
HESSJ1912+101	2.9e+36	169	4.8 +0.7 -0.5	7.27 +/- 3.731	···

Table 9.1 (cont'd)

Source	$\dot{E}$ ($\mathrm{erg\,s^{-1}}$)	τ (kyr)	Distance (kpc)	Flux TeV	Flux Xray
VERJ1930+188	1.2e+37	2.89	9.0 +2.0 -7.0	2.26 +/- 1.28	5.23 +/- 0.122
VERJ1959+208	1.6e+35	⋯	2.5 +1.0 -1.0	⋯	⋯
MGROJ2019+37	3.4e+36	17.2	8.0 +0.0 -0.0	⋯	⋯
MGROJ2228+611	2.2e+37	10.5	7.3 +2.2 -3.3	⋯	0.884 +/- 0.0206

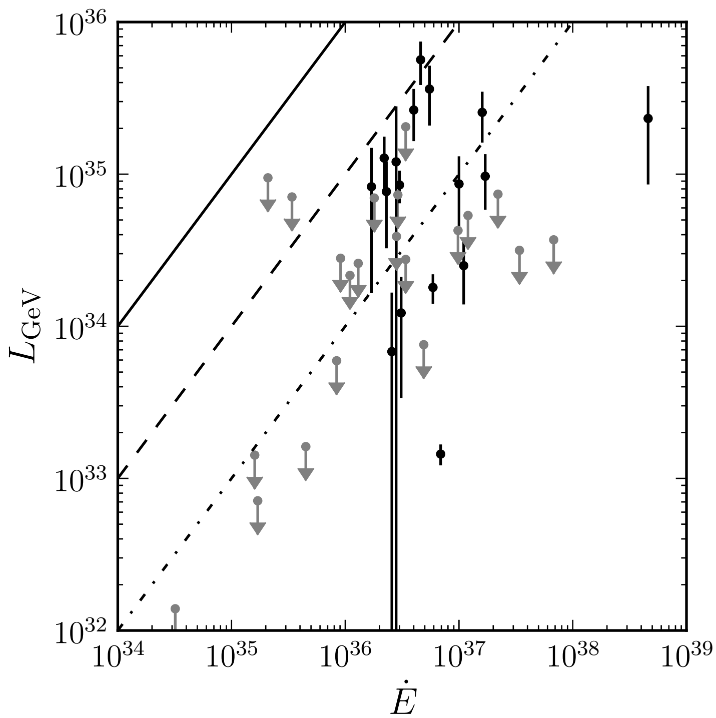

Figure 9.1 ...

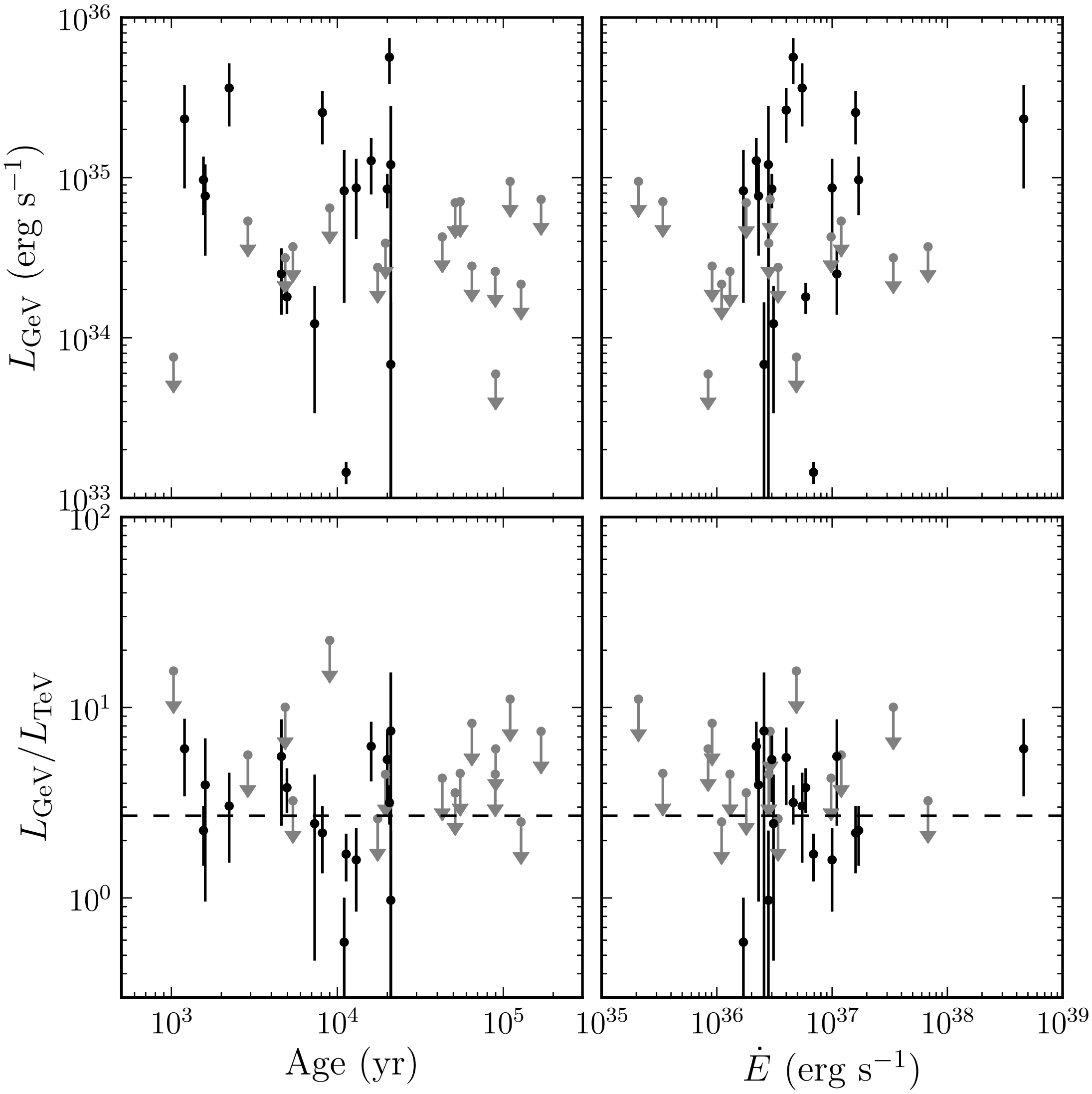
10^{36}
10^{35}
10^{34}
10^{33}
L_{GeV} (erg s^{-1})
10^{2}
10^{1}
10^{0}
$L_{\mathrm{GeV}}/L_{\mathrm{TeV}}$
10^{3}
10^{4}
10^{5}
Age (yr)
10^{35}
10^{36}
10^{37}
10^{38}
10^{39}
$\dot{E}$ (erg s^{-1})

Figure 9.2 ...

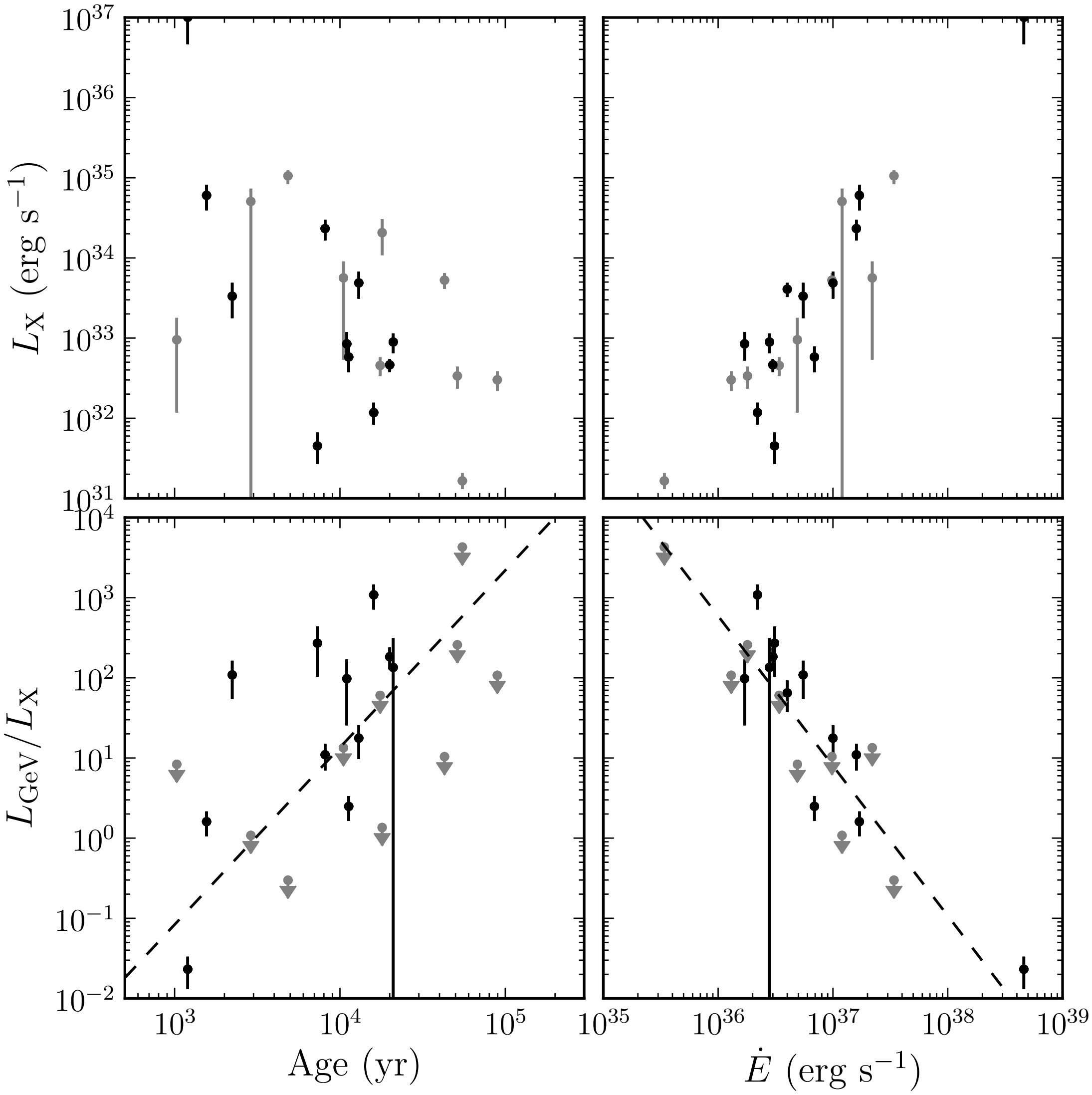

Figure 9.3 ...

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