

THE EFFECTS OF BINARITY ON PLANET OCCURRENCE RATES MEASURED BY TRANSIT SURVEYS

L. G. BOUMA,¹ K. MASUDA,¹ AND J. N. WINN¹

¹*Department of Astrophysical Sciences, Princeton University, 4 Ivy Lane, Princeton, NJ 08540, USA*

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ABSTRACT

This study develops simplified models of signal-to-noise limited transit surveys, in order to clarify the biases that stellar binarity introduces to occurrence rate measurements. Particular attention is paid to dilution of observed planet radii, errors in star counts, and detection efficiencies. Free parameters in our models include the true radius distribution of planets, and the relative numbers of planets orbiting secondaries, primaries, and single stars. In the simplest possible case, all stellar systems are either single or twin binaries, and all planets are identical. If so, then we find that ignoring binarity leads to an underestimate of the occurrence at the true radius by a multiplicative factor of 0.76 (assuming a 10% twin binary fraction). Using more realistic models for the stellar population and planetary radii, we find that the occurrence of Earth-sized planets can be overestimated by at most $\approx 50\%$ when ignoring binarity – at present far smaller than systematic uncertainties on real η_{\oplus} measurements. We quantify how gaps in the intrinsic planet radius distribution are filled in when binarity is neglected. Our models suggest that for most of the observed radius distribution ($r > 2r_{\oplus}$), ignoring binarity only leads to errors in inferred occurrence rates at the $\lesssim 10\%$ level. One corollary is that binarity is unlikely to strongly influence the HJ rate discrepancy.

Keywords: methods: data analysis — planets and satellites: detection
— surveys

1. INTRODUCTION

A group of binarity-ignoring astronomers wants to measure the mean number of planets of a certain type per star of a certain type. They perform a wide-field photometric search for planets that transit, and discover all for which

$$\left(\frac{S}{N}\right) = (\text{const}) \cdot \delta_{\text{obs}} L_{\text{sys}}^{1/2} d^{-1} > \left(\frac{S}{N}\right)_{\text{floor}}. \quad (1)$$

The observed transit depth δ_{obs} constitutes the signal, S . The survey is shot-noise dominated, and the fractional photon noise N scales as the inverse root of the photon flux received from any stellar system, $N \propto (L_{\text{sys}}/d^2)^{-1/2}$, for L_{sys} the system luminosity and d its distance. The solid angle coverage, telescope area, and survey duration are part of the “const” term of Eq. 1. The transit duration also affects detectability, but we omit it in this work for brevity.

To compute an occurrence, at every planetary and stellar property the astronomers choose the stars around which the planets of interest appeared to be searchable. Correcting for the geometric transit probability p_{tra} , they report an apparent rate density,

$$\Gamma_a(\mathcal{P}_a, \mathcal{S}_a) = \frac{n_{\text{det}}(\mathcal{P}_a, \mathcal{S}_a)}{N_{\text{s,a}}(\mathcal{P}_a, \mathcal{S}_a)} \times \frac{1}{p_{\text{tra}}(\mathcal{P}_a, \mathcal{S}_a)}. \quad (2)$$

where \mathcal{P}_a , \mathcal{S}_a are the apparent planetary and stellar parameters. The quantity $N_{\text{s,a}}(\mathcal{P}_a, \mathcal{S}_a)$ is the number of unresolved point-sources on the sky that appear to be searchable, and $n_{\text{det}}(\mathcal{P}_a, \mathcal{S}_a)$ is the number of detected planets, per unit \mathcal{P}_a and \mathcal{S}_a .

There are many potential pitfalls. Some genuine transit signals can be missed by the detection pipeline. Some apparent transit signals are spurious, from noise fluctuations, failures of ‘detrending’, or instrumental effects. Stars and planets can be misclassified due to statistical and systematic errors in the measurements of their properties. Poor angular resolution causes false positives due to blends with background eclipsing binaries. *Et cetera*.

Here we focus on problems that arise from the fact that many stars exist in multiple star systems. For simplicity, we consider only binaries, and we assume that they are all bound and spatially unresolved.

An immediate complication is that, due to dynamical stability or some aspect of planet formation, the occurrence rate of planets might differ between binary and single-star systems. In an observed sample of stars that contains both singles and binaries, the detected planets would then be drawn from different intrinsic distributions for the occurrence rates in single star systems, about primaries, and about secondaries (*e.g.*, Wang et al. 2015a).

Outside of astrophysical differences, every term in Eq. 2 is observationally-biased. There are errors in $n_{\text{det}}(\mathcal{P}_a, \mathcal{S}_a)$ due to misclassification of both planet radii and stellar properties. There are errors in $N_{\text{s,a}}(\mathcal{P}_a, \mathcal{S}_a)$ because an apparently searchable point on the sky might include two searchable stars. There are errors in $p_{\text{tra}}(\mathcal{P}_a, \mathcal{S}_a)$ because

stars in binaries may have different masses and radii than assumed in the single-star case. Binarity might also affect the observers’ ability to correctly select searchable stars (but see Sec. 5 for discussion of this issue).

Correcting for binarity’s observational biases is non-trivial. For instance, in counting the number of searchable stars, even after realizing that binaries count as two stars, one must note that the multiplicity fraction of stars that are searchable for a given observed transit depth is greater than that the multiplicity fraction in a volume limited sample. This is the familiar Malmquist bias: at a fixed δ_{obs} , binaries are searchable out to larger distances than single stars because they are more luminous. While this is what biases detectability, separate effects bias interpretation: ignoring binarity, the (apparent) radii assigned to planets in binaries will typically be smaller than the true radii. The apparent and true radii differ because of diluting flux, and possibly because the planet is assumed to orbit the wrong star.

To gain intuition for the many observational biases at play, we consider a set of idealized transit surveys. First, we discuss the simplest possible cases (Sec. 2), in which all planets have identical properties, and all stars have identical properties except that some are in binaries. The straightforward thought-experiments of Sec. 2 are then generalized in Sec. 3 to allow for varying populations of secondaries, and for varying intrinsic occurrence rate distributions. Sec. 4 investigates questions of interest, including (a) measurements of η_{\oplus} , (b) the “hot Jupiter rate discrepancy”, and (c) [Fulton et al. \(2017\)](#)’s recently discovered “valley”. We discuss what our models suggest about binarity’s biasing effects in Sec. 5, and conclude in Sec. 6.

2. THE SIMPLEST POSSIBLE TRANSIT SURVEY MODELS

2.1. *Twin binaries, identical stars, identical planets*

Since the effects of binarity are most pronounced when the two components are similar, we begin by considering a universe in which all planets are identical, and all stars are identical except that some fraction of them exist in binaries. The goal of this thought experiment is to compare the occurrence inferred by observers ignoring binarity with the true occurrence in single star systems.

First, what are the possible apparent radii of detected planets? If the true planet radius is $r = r_p$, then planets in singles will be detected with the identical apparent radii: $r_a = r_p$. Conversely, all detections in twin binaries will have apparent radii $r_a = r_p/\sqrt{2}$, because they are diluted. For an unresolved point-source,

$$\delta_{\text{obs}} = \left(\frac{r_a}{R_a} \right)^2 = \left[\frac{r}{R_{\text{host}}} \right]^2 \times \frac{L_{\text{host}}}{L_{\text{sys}}(M_1, q)}, \quad (3)$$

where R_{host} (L_{host}) is the planet-host’s stellar radius (luminosity), R_a is the apparent stellar radius, $L_{\text{sys}} = L_1 + L_2 = L(M_1) + L(qM_1)$ is the system luminosity, and $q = M_2/M_1$ is the binary mass ratio. We assume that the stellar radius and luminosity are uniquely related to the stellar mass. The observers, with no a priori knowledge of

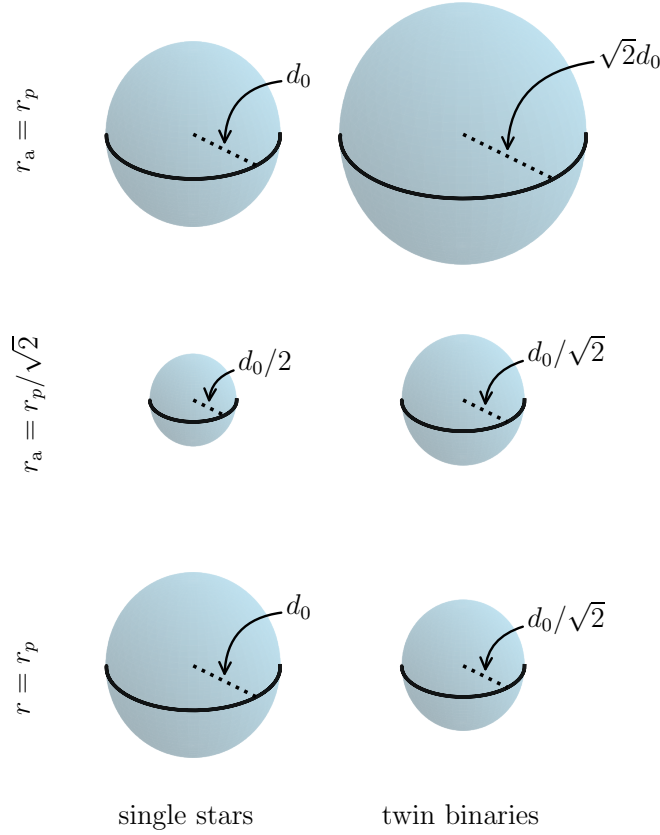


Figure 1. Cartoon of the searchable volumes for single stars (*left*), and twin binaries (*right*). This model assumes all stars have equal mass and luminosity, and all planets have the same true radius $r = r_p$. *Top:* At an apparent radius $r_a = r_p$, the observer searched twin binaries out to a larger distance, $\sqrt{2}\times$ that of single stars. Because of dilution, there are no planets in twin binaries with $r_a = r_p$. *Middle:* At an apparent radius $r_a = r_p/\sqrt{2}$, the searched volumes are half those at $r_a = r_p$. The only detected planets with $r_a = r_p/\sqrt{2}$ orbit twin binaries. *Bottom:* Planets with a true radius $r = r_p$ are searchable to a maximum distance d_0 around singles, and $d_0/\sqrt{2}$ around twin binaries.

the distance to a system, would measure a flux, assume a stellar mass $M_a = M_1$, and estimate both a stellar radius $R_a = R(M_a)$ and an under-luminous system luminosity of $L_{\text{sys},a}(M_a) = L(M_a)$.

The planets will be detected if and only if their host stars are searchable. For a fixed signal-to-noise floor, there is a maximum searchable distance (Pepper et al. 2003; Pepper & Gaudi 2005). From Eq. 1, this maximum searchable distance scales as

$$d(\delta_{\text{obs}}, L_{\text{sys}}) \propto \delta_{\text{obs}} \cdot L_{\text{sys}}^{1/2}. \quad (4)$$

Assuming that stars are uniformly distributed in space, the number of searchable stars N_s is then proportional to

$$N_s(\delta_{\text{obs}}, L_{\text{sys}}) \propto n \delta_{\text{obs}}^3 L_{\text{sys}}^{3/2}, \quad (5)$$

where n is the number per unit volume of *e.g.*, single star, or binary systems. We neglect the dependence of n on stellar type.

The searchable volumes for our thought experiment are illustrated in Fig. 1. At an apparent radius equal to the true radius ($r_a = r_p$), single stars with $d < d_0$ are searchable. Binaries with $r_a = r_p$ are also searchable, out to $\sqrt{2}d_0$. However, in this model no such systems exist; all planets in twin binaries have true radii $r = r_p$, and so are observed with apparent radii $r_a = r_p/\sqrt{2}$, out to $d_0/\sqrt{2}$. At an apparent radius of $r_a = r_p/\sqrt{2}$, there could also be planets with true radii $r = r_p/\sqrt{2}$ orbiting singles — however none of these exist in this model.

How many planets are detected? If we assume that there are $N_s^0(r_a)$ singles that are searchable for planets with r_a , and that there are Z_0 planets per single, then single stars contribute

$$n_{\text{det}}^0(r_a) = N_s^0(r_a) Z_0 p_{\text{tra}} \cdot \delta(r_a - r_p) \quad (6)$$

planet detections (per unit r_a), where δ is the Dirac delta function.

We can write similar equations for the binaries. First though, we define a ratio $\mu \equiv N_s^b(r_a)/N_s^0(r_a)$ between the number of searchable binary systems and the number of searchable singles at a given r_a . This quantity is useful because although N_s^b and N_s^0 vary as a function of the observed transit depth, μ remains fixed. Applying this definition, at $r_a = r_p/\sqrt{2}$, some detections will come from primaries,

$$n_{\text{det}}^1(r_a) = \mu N_s^0(r_a) Z_1 p_{\text{tra}} \cdot \delta\left(r_a - \frac{r_p}{\sqrt{2}}\right), \quad (7)$$

and some from secondaries,

$$n_{\text{det}}^2(r_a) = \mu N_s^0(r_a) Z_2 p_{\text{tra}} \cdot \delta\left(r_a - \frac{r_p}{\sqrt{2}}\right), \quad (8)$$

where Z_1 is the number of planets per primary, Z_2 is the number of planets per secondary, and p_{tra} is the same for identical stars. In the case of twin binaries, Eq. 5 gives $\mu = 2^{3/2}n_b/n_s = 2^{3/2}\text{BF}/(1 - \text{BF})$, for BF the binary fraction¹.

The last item we need to write down the apparent rate density $\Gamma_a(r_a)$ is the number of unresolved point-sources on sky, $N_{s,a}$. Assuming that the apparent stellar radius is fixed, then given the apparent planet radius r_a ,

$$N_{s,a}(r_a) = N_s^0(r_a) \times (1 + \mu) \quad (9)$$

relates the number of apparently searchable point-sources to the number of searchable singles. From the discussion above, it should be clear that there are in fact $N_s^0(r_a) \times (1 + 2\mu)$ searchable stars at any given r_a ; the observers are under-counting.

Using Eq. 2, the apparent rate density computed without considering binarity is

$$\Gamma_a(r_a) = \frac{1}{1 + \mu} Z_0 \cdot \delta(r_a - r_p) + \frac{\mu}{1 + \mu} (Z_1 + Z_2) \cdot \delta\left(r_a - \frac{r_p}{\sqrt{2}}\right). \quad (10)$$

¹ The binary fraction is the fraction of systems in a volume-limited sample that are binary. It is equivalent to the multiplicity fraction if there are no triple, quadruple, or higher order multiples.

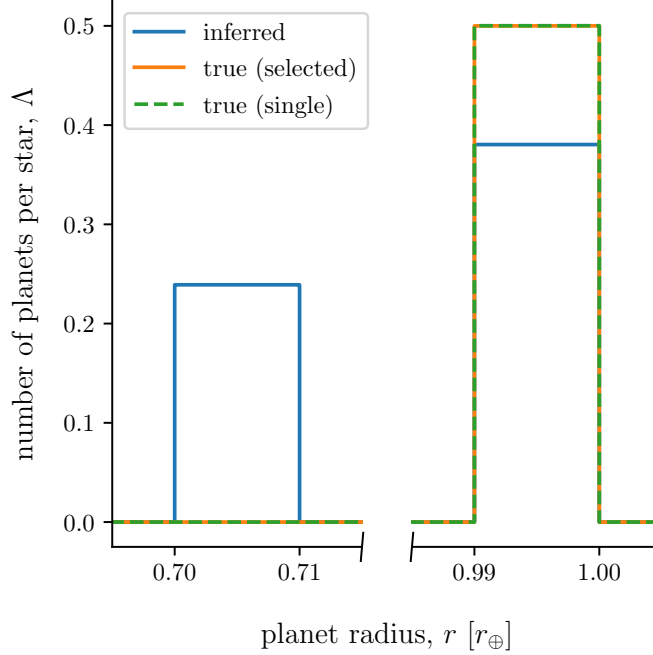


Figure 2. Inferred occurrence rates over $0.01r_\oplus$ bins in planet radius, for a model with identical stars, identical planets, and twin binaries. We assume a twin binary fraction of $\text{BF} = 0.1$, and that all stars (single, primary, and secondary) host planets at equal rates. If the true planet radius is r_p , all planets detected in binaries will have apparent radii $r_a = r_p/\sqrt{2}$.

The number of searchable singles at a given apparent radius, $N_s^0(r_a)$, makes no appearance. The observer’s errors in Eq. 10 are (A) under-counting the number of searchable stars at any apparent radius; and (B) inferring radii in binary systems that are all $\sqrt{2}$ too small.

To assess the severity of these errors, we need to assume a binary fraction, and to know the true number of planets per single, primary, and secondary star (Z_0, Z_1 , and Z_2). For the former, [Raghavan et al. \(2010\)](#) found a multiplicity fraction of 0.44 for primaries with masses from $0.7M_\odot$ to $1.3M_\odot$, which we take as a binary fraction. “Twin binaries” are perhaps 10-20% of the binaries in a volume-limited sample, depending on how “twin” is defined. Thus a representative twin binary fraction for this model is $\text{BF} \approx 0.05\text{-}0.1$. For the latter, we take $Z_0 = Z_1 = Z_2 = 0.5$, and plot the resulting occurrence rates as a function of planet radius in Fig. 2.

Our point for comparison is the rate density for single stars,

$$\Gamma_0(r) = Z_0 \cdot \delta(r - r_p). \quad (11)$$

We can define a rate density “correction factor”, X_Γ , as the ratio of the apparent to single star rate density,

$$X_\Gamma(r) \equiv \left. \frac{\Gamma_a(r_a)}{\Gamma_0(r)} \right|_{r_a \rightarrow r}. \quad (12)$$

Continuing in the assumption that the number of planets per single, primary, and secondary star are equal, we find a rate density correction factor at the true planet radius of $1/(1+\mu)$. This yields a correction of 0.76 if $\text{BF} = 0.1$, and 0.87 if $\text{BF} = 0.05$. Rephrasing the result, if the twin binary fraction were $(2^{3/2} + 1)^{-1} \approx 0.26$, then the apparent rate at $r_a = r_p$ would be half the true rate at $r = r_p$. For most stellar types, the twin binary fraction is not that high.

An alternative assumption is that secondaries do not host planets. In that case, $Z_0 = Z_1$, and $Z_2 = 0$. The correction to the rate density at the true planet radius becomes $(1+2\mu)/(1+\mu)^2$. This evaluates to 0.94 if $\text{BF} = 0.1$, and 0.98 if $\text{BF} = 0.05$. While it seems unlikely that secondaries are planet-less, it is worth noting that Γ_a/Γ is sensitive to the relative number of planets per single, primary, and secondary.

2.2. Twin binaries, identical stars, two planet sizes

A simple extension to the previous example will help us further distinguish the detected populations. Consider now a universe that is the same in every respect to that of Sec. 2.1, except now half of planets have radii r_p , while the other half have $r_p/\sqrt{2}$. The true rate densities are

$$\Gamma_i(r) = \frac{Z_i}{2} \left[\delta(r - r_p) + \delta\left(r - \frac{r_p}{\sqrt{2}}\right) \right], \quad \text{for } i \in \{0, 1, 2\}, \quad (13)$$

where as always $i = 0$ corresponds to singles, $i = 1$ to primaries, and $i = 2$ to secondaries, and the Z_i 's are the number of planets per star of each type.

Following an identical line of reasoning as before, the apparent rate density can be written

$$\begin{aligned} \Gamma_a(r_a) = \frac{1}{2(1+\mu)} & \left[Z_0 \cdot \delta(r_a - r_p) + [Z_0 + \mu(Z_1 + Z_2)] \cdot \delta\left(r_a - \frac{r_p}{\sqrt{2}}\right) \right. \\ & \left. + \mu(Z_1 + Z_2) \cdot \delta\left(r_a - \frac{r_p}{2}\right) \right]. \end{aligned} \quad (14)$$

The important qualitative point of this equation is that at $r_a = r_p/\sqrt{2}$, two populations contribute. The first population is singles with $r = r_p/\sqrt{2}$, and the second is twin binaries with $r = r_p$. Both populations are detected out to the same maximum searchable distance. Just as with Eq. 9, we will have a cancellation of $N_s^0(r_a)$ terms, leaving only the μ weights.

3. GENERAL FORMULA FOR APPARENT OCCURRENCE RATE

To generalize the procedure of Sec. 2, we consider a S/N limited transit survey in which all the primaries and single stars have identical properties. For example, they might all be solar analogs, roughly similar to the stars that *Kepler* surveyed (Batalha et al. 2010). We let the properties of the secondaries vary, and assume that there are some functions $L(M)$ and $R(M)$ that specify a star's luminosity and radius in terms of its mass. We then allow for an arbitrary volume-limited mass ratio distribution

$f(q)$, and also for arbitrary true rate densities for singles, primaries, and doubles ($\Gamma_0, \Gamma_1, \Gamma_2$). Since the properties of the secondaries vary, we must consider not only the dependence of the rate densities on planet radius, but also on stellar mass. We presume that the observers always take the properties of a binary system to be those of the primary.

Given this laundry-list of assumptions, a general formula for the apparent rate density follows:

$$\Gamma_a(r_a, M_a) = \frac{1}{1 + \mu(\text{BF}, M_a)} \times \left\{ \Gamma_0(r_a, M_a) + \frac{\text{BF}}{1 - \text{BF}} \left[\int \frac{dq}{\mathcal{A}^4} f(q) \Gamma_1\left(\frac{r_a}{\mathcal{A}}, M_a\right) + \int \frac{q dq}{\mathcal{A}^3 \mathcal{B}} f(q) \Gamma_2\left(\frac{r_a}{\mathcal{B}}, q M_a\right) \frac{R(q M_a)}{R(M_a)} q^{-1/3} \right] \right\}, \quad (15)$$

where

$$\mathcal{A}(q, M_a) = \sqrt{\frac{L(M_a)}{L_{\text{sys}}(M_a, q)}} \quad (16)$$

and

$$\mathcal{B}(q, M_a) = \frac{R(M_a)}{R(q M_a)} \sqrt{\frac{L(q M_a)}{L_{\text{sys}}(M_a, q)}} \quad (17)$$

are the dilution terms appropriate to primaries and secondaries, and as before BF is the volume-limited binary fraction.

A derivation of Eq. 15 is given in the [Appendix](#); here we highlight its qualitative features. The terms inside the curly braces $\{\dots\}$ are the sum of the apparent rate densities from singles, primaries, and secondaries. The contribution from singles is not affected by binarity. In the contributions from primaries and secondaries, planets with (r_a, M_a) are associated with systems of many different planetary and stellar properties. This means that $n_{\text{det}}(r_a, M_a)$ is given by a convolution of the true rate density, $\Gamma(r, M)$, and the number per unit (r_a, M_a) of searchable stars that give (r_a, M_a) when the true system actually has (r, M) . This convolution has been evaluated, leaving only a dependence on the mass ratio distribution. Finally, the latter term in the apparent rate density of secondaries comes from a correction for the overestimated transit probability.

Ratio of searchable binaries to singles—We must also specify the exact form of $\mu(\text{BF}, M_a)$, the ratio of the number of searchable binary systems to singles. Given the observed signal δ_{obs} and apparent stellar mass M_a , recall that the maximum searchable distance for singles and binaries are proportional to $\delta_{\text{obs}} \cdot L(M_a)^{1/2}$ and $\delta_{\text{obs}} \cdot [L(M_a) + L(q M_a)]^{1/2}$. Applying Eq. 5,

$$\mu(\text{BF}, M_a) = \frac{N_s^b(r_a)}{N_s^0(r_a)} = \int_0^1 \frac{n_b}{n_s} \left[1 + \frac{L(q M_a)}{L(M_a)} \right]^{3/2} f(q) dq. \quad (18)$$

If BF is independent of q , $L \propto M^\alpha$, and $f(q) \propto q^\beta$, this reduces to

$$\mu = \frac{\text{BF}}{1 - \text{BF}} \cdot \frac{1}{1 + \beta} \int_0^1 (1 + q^\alpha)^{3/2} q^\beta dq, \quad (19)$$

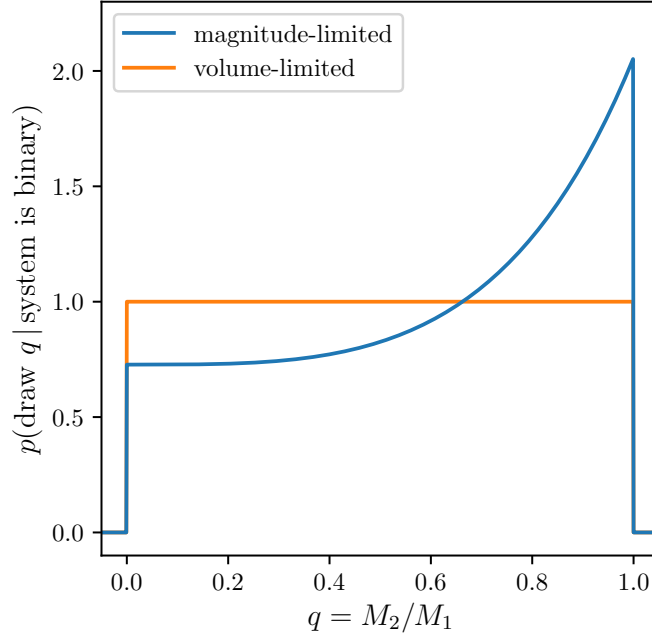


Figure 3. The mass ratio distribution for a magnitude-limited sample of binary stars, in which the underlying volume-limited distribution is uniform, qualitatively similar to *e.g.*, [Raghavan et al. \(2010\)](#)’s Fig. 16. At a given observed depth, the searchable binaries in a transit survey are magnitude-limited.

which may be written in a closed form using the hypergeometric function.

This can be understood as a Malmquist bias: at fixed observed transit depth, for singles and primaries with the same luminosity, binaries are searchable to a greater distance. In particular, the sample of searchable binaries is magnitude-limited. Given a searchable binary, the probability it having a mass ratio q scales as

$$p(\text{draw } q \mid \text{system is binary}) \propto (1 + q^\alpha)^{3/2} q^\beta \quad (20)$$

where q^β is the volume-limited probability of drawing a binary of mass ratio q , and the first term is the Malmquist bias. We show this magnitude-limited mass ratio distribution for the $\beta = 0$ case in Fig. 3. In Monte Carlo simulations of transit surveys, it is important to draw binaries from a correctly biased mass-ratio distribution (*e.g.*, [Bakos et al. 2013](#); [Sullivan et al. 2015](#); [Günther et al. 2017](#)).

4. MORE COMPLICATED TRANSIT SURVEY MODELS

The following section applies our general equation for the apparent rate density (Eq. 15) to study binarity’s effects in more realistic regimes. In general, we will write the rate density for each type of star, $\Gamma_i(r)$, as the product of a shape function and a constant:

$$\Gamma_i(r) = \frac{d\Lambda_i}{dr} = Z_i f_i(r), \quad \text{for } i \in \{0, 1, 2\}, \quad (21)$$

where as awlays, $i = 0$ corresponds to single stars, $i = 1$ to primaries of binaries, and $i = 2$ to secondaries of binaries. The shape function is normalized to unity. The

Z_i 's are each system type's occurrence rate Λ_i , integrated over all planetary radii. In other words, they are number of planets per single, primary, or secondary star. This section will consider the effects of varying both $f_i(r)$ and also the relative values of the Z_i 's.

4.1. Power law planet radius distributions

4.1.1. Twin binaries

We begin introducing realism by keeping all binaries as twins, but letting the radius distribution of planets vary. Specifically, we take

$$\Gamma_i(r) = Z_i f(r) = Z_i r^\delta / \mathcal{N}_r, \quad (22)$$

for $f(r)$ the radius shape function, and \mathcal{N}_r the shape function's normalization. The resulting apparent rate density is quite similar to that of the twin binary, fixed-planet case (Eq. 10), except that the radius dependence and normalization are slightly different:

$$\Gamma_a(r_a) = \frac{r_a^\delta}{\mathcal{N}_r} \left[\frac{Z_0}{1 + \mu} + 2^{\frac{\delta+1}{2}} \frac{\mu}{1 + \mu} (Z_1 + Z_2) \right]. \quad (23)$$

As in Sec. 2.1, $\mu = 2^{3/2} \text{BF} / (1 - \text{BF})$. If the Z_i 's are equal, the ‘‘correction factor’’ relative to the rate density for singles then becomes quite simple:

$$\left. \frac{\Gamma_a(r_a)}{\Gamma_0(r)} \right|_{r_a \rightarrow r} = \frac{1 + 2^{\frac{\delta+3}{2}} \mu}{1 + \mu}. \quad (24)$$

For the case of $\text{BF} = 0.1$, $\mu \approx 0.153$. Taking $\delta = -2.92$ from Howard et al. (2012), we get a correction factor $\Gamma_a/\Gamma_0 = 1.004$. In other words, the apparent rate density is an *overestimate* compared to the rate density of single stars, with a relative difference $\delta\Gamma_0 = |\Gamma_0 - \Gamma_a|/\Gamma_0$ of 0.4%. This is quite a small effect! Note though that it could change if the Z_i 's are not equal. Also, a power law radius distribution $f(r) \propto r^\delta$ diverges for $\delta < 0$; Howard et al. (2012)'s results indicate that this approximation is good for $r_a \gtrsim 2r_\oplus$. Similarly, it is nonsensical for $f(r)$ to be finite above some upper radius limit r_u , perhaps around $\approx 24r_\oplus$, based on the most inflated hot Jupiters. An upper cut-off of r_u would lead to smaller values of $\Gamma(r_a)$ down to $r_u/\sqrt{2}$, compared to an intrinsic power law without the cut-off. We are not particularly interested in this effect; there are better models for hot Jupiter occurrence rates that we will discuss in Sec. 4.3. If the assumptions behind this model (listed in Sec. 3) are at all applicable in real transit surveys, this suggests that for $2r_\oplus \lesssim r_a \lesssim 17r_\oplus$, binarity's impact on derived occurrence rates is quite small.

4.1.2. Varying binaries

To confirm the intuition that non-twin binaries do not affect the above result, we now assume $f(q) = q^\beta / \mathcal{N}_q$, for \mathcal{N}_q the normalization. This changes μ (Eq. 18 becomes Eq. 19). It may also affect the rate density, which could be a function of the stellar

mass, which is now varying for secondaries (we assume that singles and primaries have the same mass). Absorbing this dependence into a power law as well,

$$\Gamma_i(r, M) = Z_i \times \frac{r^\delta}{\mathcal{N}_r} \times \frac{M^\gamma}{\mathcal{N}_M}, \quad (25)$$

where Z_i is dimensionless, and the normalization constants carry the units (each side has units $[r^{-1}M^{-1}]$). We assume that stars are a one-parameter family, given by $R \propto M \propto L^{\frac{1}{\alpha}}$, so that a given value of q determines everything about a secondary.

Using Eq. 15, we can show that when $Z_0 = Z_1 = Z_2$,

$$X_\Gamma \equiv \frac{\Gamma_a(r_a, M_a)}{\Gamma_0(r, M)} \Big|_{r_a \rightarrow r, M_a \rightarrow M} \quad (26)$$

$$= \frac{1}{1 + \mu} \left[1 + \frac{1}{\mathcal{N}_q} \frac{\text{BF}}{1 - \text{BF}} \times \left(\int dq q^\beta (1 + q^\alpha)^{\frac{\delta+4}{2}} + \int dq q^{\beta+\gamma+\delta+\frac{8}{3}} (1 + q^\alpha)^{\frac{3}{2}} (1 + q^{-\alpha})^{\frac{\delta+1}{2}} \right) \right]. \quad (27)$$

For $\alpha = 3.5$, $\beta = 0$, $\gamma = 0$, $\delta = -2.92$, the summed integrals in Eq. 27 give $(\dots) \approx 1.41425$. For $\text{BF} = 0.44$, this yields $\Gamma_a/\Gamma_0 = 1.014$; the apparent rate density is higher than the rate density around singles by 1.4%. To summarize, this model suggests that for a power-law radius distribution ($2r_\oplus \lesssim r_a \lesssim 17r_\oplus$) in which there are the same number of planets orbiting singles, primaries, and secondaries, binarity influences apparent planet occurrence rates around Sun-like stars at the \sim few percent level.

4.2. Fixed primaries, varying secondaries, broken power-law planets

Though the details of the radius distribution $\Gamma_i(r)$ at $r < 2r_\oplus$ are still an active topic of research, we can consider how binarity effects this regime under various reasonable assumptions. For instance, we can assume that the true radius shape function is

$$f(r) \propto \begin{cases} r^\delta & \text{for } r \geq 2r_\oplus \\ \text{constant} & \text{for } r \leq 2r_\oplus. \end{cases} \quad (28)$$

As before, all single and primary stars in this model have identical properties. The secondaries have masses, radii, and luminosities that vary between systems: $R \propto M \propto L^{\frac{1}{\alpha}}$. Our “nominal model” remains the same: $\alpha = 3.5$, $\beta = \gamma = 0$, $\delta = -2.92$.

At apparent radii $r_a > 2r_\oplus$, the results of this model are the same as those from Sec. 4.1.2. For $r_a < 2r_\oplus$, the equations are quite tedious, though still tractable. To compute $\Gamma_a(r_a)$, we simply insert Eq. 28 into Eq. 15, and integrate using a computer program. We refer the interested reader to our online implementation². The output is validated against analytic predictions in the $r_a > 2r_\oplus$ and the $r_a < 2r_\oplus/\sqrt{2}$ regimes, and is plotted in Fig. 4.

² https://github.com/lgbouma/binary_biases

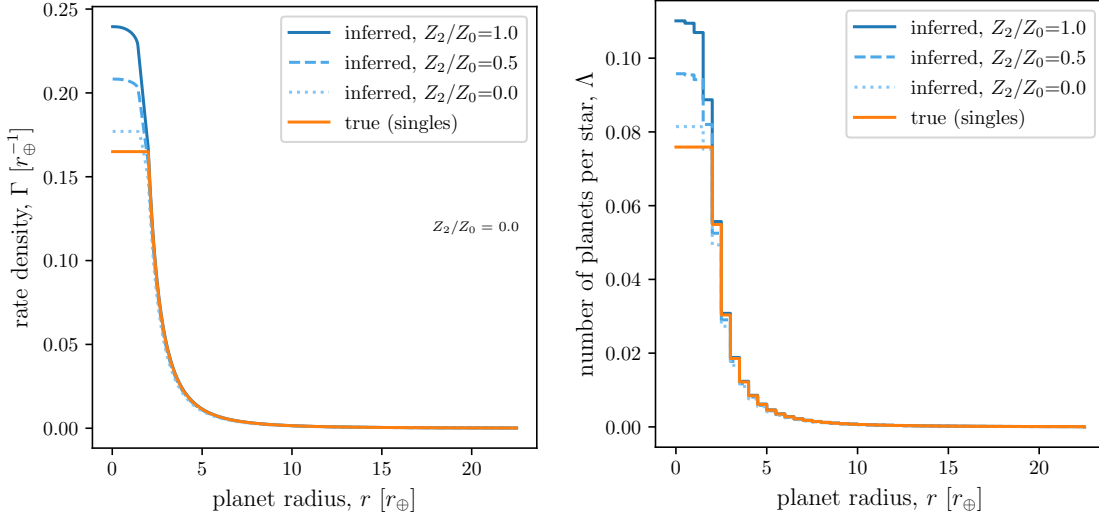


Figure 4. *Left:* rate density and *right:* rate (over $0.5r_{\oplus}$ bins) as a function of planet radius for the distribution specified by Eq. 28. This model has identical primaries and single stars, and varying secondaries. We use an arbitrary normalization of $Z_0 = 0.5$ throughout.

The rate of Earth analogs—The most immediately arresting result there is a “bump” in the apparent rate density at $r_a < 2r_{\oplus}$: the true rate for singles is less than the inferred rate. For the case in which secondaries host as many (half as many) planets as single stars, this means an overestimate of the absolute occurrence rate by $\approx 50\%$ ($\approx 25\%$). The “bump” exists even for the $Z_2/Z_0 = 0$ case as a $\approx 10\%$ effect.

Hot Jupiter occurrence rates—Taking Fig. 4 and integrating the rate density, we can compare the apparent hot Jupiter occurrence rate with the true rate for singles. In the most extreme case of $Z_2/Z_0 = 0$, we find that $\Lambda_{\text{HJ},0}/\Lambda_{\text{HJ},a} = 1.13$, where

$$\Lambda_{\text{HJ},a} = \int_{8r_{\oplus}}^{\infty} \Gamma_a(r) dr, \quad (29)$$

and similarly for $\Lambda_{\text{HJ},0}$. If $Z_2/Z_0 > 0$, binarity affects this value less: when $Z_2/Z_0 = 0.5$, we find $\Lambda_{\text{HJ},0}/\Lambda_{\text{HJ},a} = 1.06$

4.3. Further models: radius gaps, Gaussian HJ distributions

The radius distribution specified by Eq. 28 misses some important features recently reported by state of the art occurrence rate studies.

Precise features of the radius valley—In particular, [Fulton et al. \(2017\)](#) recently reported a “gap” in the radius-period plane ([Petigura et al. 2017a](#); [Johnson et al. 2017](#)). The existence of the gap has been independently corroborated from a sample of KOIs with asteroseismically-determined stellar parameters ([Van Eylen et al. 2017](#)). Precise measurement of the gap’s features, in particular its width, depth, and shape, will require accurate occurrence rates. To illustrate binarity’s role in this problem, we make identical assumptions as in Sec. 4.2, but instead assume an intrinsic radius

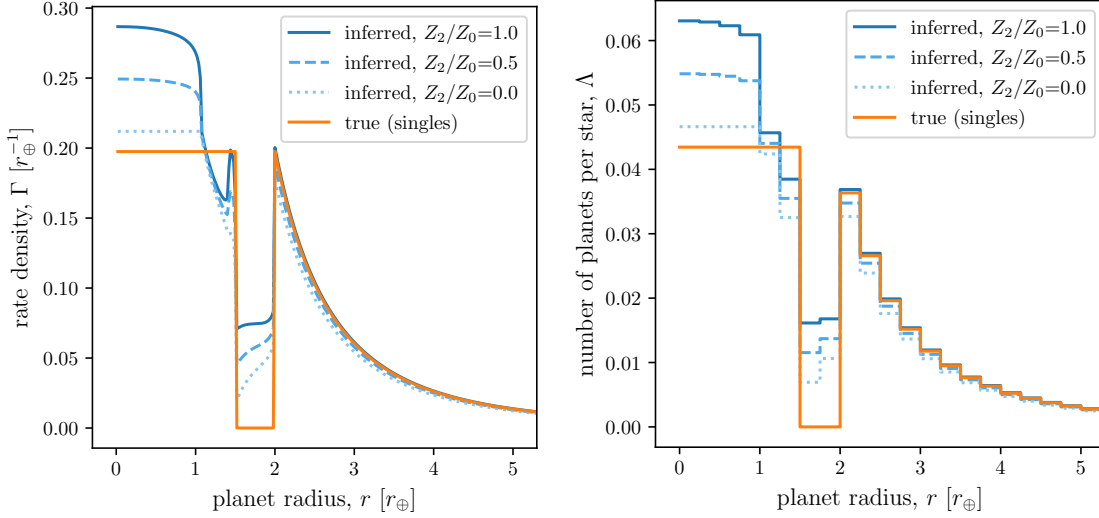


Figure 5. *Left:* rate density and *right:* rate (over $0.25r_{\oplus}$ bins) as a function of planet radius in a model with a radius gap (Eq. 30). Similar to Fig. 4, this model has fixed primaries and single stars, but varying secondaries.

distribution

$$f(r) \propto \begin{cases} r^{\delta} & \text{for } r \geq 2r_{\oplus}, \\ 0 & \text{for } 1.5r_{\oplus} < r < 2r_{\oplus}, \\ \text{constant} & \text{for } r \leq 1.5r_{\oplus}. \end{cases} \quad (30)$$

The resulting true and inferred rates are shown in Fig. 5. If left uncorrected, binarity makes the gap appear more shallow, and flattens the step-function edges. Of course, other effects could also “blur” the gap in the planet radius dimension. In particular, the valley’s period-dependence is almost certainly not flat (Van Eylen et al. 2017; Owen & Wu 2017). This means that any study measuring the gap’s width or depth in the face of binarity must either perform tests at fixed orbital period, or else marginalize over period and account for the associated blurring in the planet radius dimension.

Alternative models for the HJ distribution—In the recent work by Petigura et al. (2017b), hot Jupiters appear as an island in period-radius space, rather than as a continuous component of a power law distribution. Consider instead a Gaussian radius shape function,

$$f(r) \propto \exp\left(-\frac{(r - \bar{r})^2}{2\sigma^2}\right), \quad (31)$$

with $\bar{r} = 14r_{\oplus}$ and $\sigma = 2r_{\oplus}$. As always, $\Gamma_i(r) = Z_i f(r)$. We then compute $\Gamma_a(r_a)$, and plot it in Fig. 6. Integrating to find hot Jupiter rates, we find the opposite effect as in the power law model. For $Z_2/Z_0 > 0$, the apparent HJ rate is *greater* than the true rate for singles. The effect is maximal when there are as many hot Jupiters

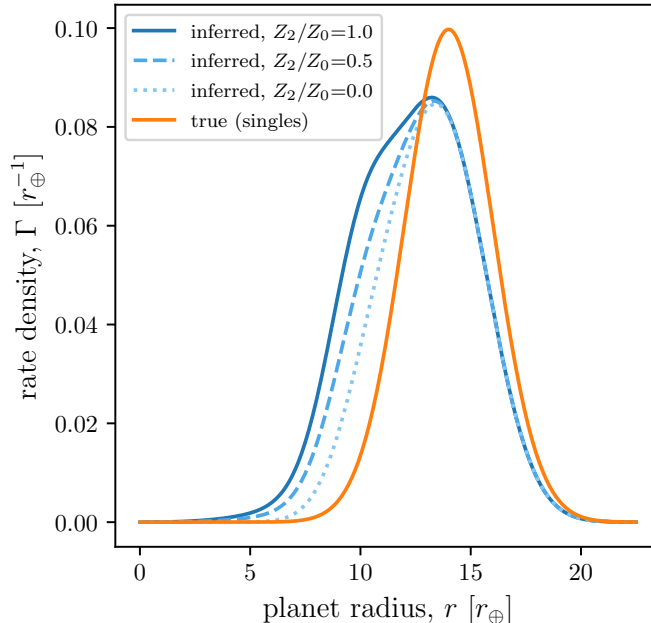


Figure 6. Rate density for a population of planets with true radii r drawn from $\mathcal{N}(14r_{\oplus}, 2r_{\oplus})$. This is similar to the hot Jupiter distribution presented by [Petigura et al. \(2017b\)](#).

orbiting secondaries as singles, in which case $\Lambda_{\text{HJ},0}/\Lambda_{\text{HJ},a} = 0.84$. For the case when no hot Jupiters orbit secondaries, $\Lambda_{\text{HJ},0} \approx \Lambda_{\text{HJ},a}$ to within one percent.

5. DISCUSSION

How has binarity been considered in occurrence rate measurements?—This study has shown that under a reasonable set of simplifying assumptions, binarity can introduce systematic errors to star and planet counts in transit surveys, which then bias derived occurrence rates. Thus far, real-world occurrence rate calculations³ using transit survey data have mostly ignored stellar multiplicity (*e.g.*, [Howard et al. 2012](#); [Fressin et al. 2013](#); [Foreman-Mackey et al. 2014](#); [Dressing & Charbonneau 2015](#); [Burke et al. 2015](#)). For *Kepler* occurrence rates specifically, it seems that no one has yet carefully assessed binarity’s importance, or lack thereof. While we do not resolve the problem, we do suggest the approximate scale of the necessary corrections in a survey-independent manner.

Of course, on a system-by-system level stellar multiplicity affects the interpretation of planet candidates. High resolution imaging campaigns have measured the multiplicity of almost all *Kepler* Objects of Interest ([Howell et al. 2011](#); [Adams et al. 2012, 2013](#); [Horch et al. 2012, 2014](#); [Lillo-Box et al. 2012, 2014](#); [Dressing et al. 2014](#); [Law et al. 2014](#); [Cartier et al. 2015](#); [Everett et al. 2015](#); [Gilliland et al. 2015](#); [Wang et al. 2015b,c](#); [Baranec et al. 2016](#); [Ziegler et al. 2017](#)). The results of these programs have

³ A list of occurrence rate papers is maintained at https://exoplanetarchive.ipac.caltech.edu/docs/occurrence_rate_papers.html

been collected by [Furlan et al. \(2017\)](#), and they represent an important advance in understanding the KOI sample’s multiplicity statistics. In particular, they can be immediately applied to rectify binarity’s effects on the mass-radius diagram ([Furlan & Howell 2017](#)).

The high resolution imaging campaign is also beginning to connect with occurrence rate calculations. The most recent rate studies have used [Furlan et al. \(2017\)](#)’s catalog to test the effects of removing KOI hosts with known companions, which helps reduce contamination in the “numerator” of the occurrence rate ([Fulton et al. 2017](#); [Petigura et al. 2017b](#)). However, without an understanding of the multiplicity statistics of the non-KOI stars, the true number of searchable stars, and thus the true occurrence rates, will remain biased. The first-order correction that we suggest, given the impracticality of performing high-resolution imaging of every star searched by a transit survey, is to model the detection pipeline’s efficiency while accounting for binarity. For *Kepler*, this would require high resolution imaging of a comparison sample of non-KOI host stars. If the associated multiplicity statistics are then included in a model of the pipeline’s detection efficiency, and the number of searchable stars is appropriately counted (at a given δ_{obs}), it would correct many of binarity’s biases.

The rate of Earth analogs—Per *Kepler*’s primary science objective, the rate of Earth-like planets orbiting Sun-like stars has been independently measured by [Youdin \(2011\)](#); [Petigura et al. \(2013\)](#); [Dong & Zhu \(2013\)](#); [Foreman-Mackey et al. \(2014\)](#), and [Burke et al. \(2015\)](#). These efforts have found that the one-year terrestrial planet occurrence rate varies between ≈ 0.03 and ≈ 1 per Sun-like star, depending on assumptions that are made when retrieving the rate ([Burke et al. 2015](#)’s Fig. 17). In the power law model of Sec. 4.2, the inferred rate can be up to 50% higher than the true rate around single stars. Though this bias seems large, it is currently small compared to the other systematic factors that dominate the dispersion in η_{\oplus} measurements. If a future analyses determine absolute values of η_{\oplus} to better than a factor of two, binarity might merit closer attention.

One relevant caveat to our assessment of binarity’s importance for η_{\oplus} measurements is that none of our models included the rate density’s period-dependence. However, close binaries usually provoke dynamical instabilities, leading to fewer long-period planets per star (*e.g.*, [Holman & Wiegert 1999](#); [Wang et al. 2014](#); [Kraus et al. 2016](#)). This effect might further bias transit survey measurements of η_{\oplus} beyond our rough estimate.

The hot Jupiter rate discrepancy—There is at least one context in which ignoring binarity has been suggested as a possible source of discrepant occurrence rate measurements. Hot Jupiter occurrence rates measured by transit surveys ($\approx 0.5\%$) are marginally lower than those found by radial velocity surveys ($\approx 1\%$; see Table 1). Though the discrepancy has weak statistical significance ($< 3\sigma$), one reason to expect a difference is that the corresponding stellar populations have distinct metallicities.

As argued by Gould et al. (2006), the RV sample is biased towards metal-rich stars, which have been measured by RV surveys to preferentially host more giant planets (Santos et al. 2004; Fischer & Valenti 2005). Investigating the discrepancy from the metallicity angle, Guo et al. (2017) measured the *Kepler* field’s mean metallicity to be $[M/H]_{\text{Kepler}} = -0.045 \pm 0.009$, which is lower than the California Planet Search’s mean of $[M/H]_{\text{CPS}} = -0.005 \pm 0.006$. The former value agrees with that measured by Dong et al. (2014). Refitting for the metallicity exponent in $\Lambda_{\text{HJ}} \propto 10^{\beta[M/H]}$, Guo et al. found $\beta = 2.1 \pm 0.7$, and noted that this would imply that the metallicity difference could account for a $\approx 20\%$ relative difference in the measured rates between the CKS and *Kepler* samples – not a factor of two⁴. Guo et al. concluded that “other factors, such as binary contamination and imperfect stellar properties” must also be at play.

Aside from surveying stars of varying metallicities, radial velocity and transit surveys differ in how they treat binarity. Radial velocity surveys typically reject both visual and spectroscopic binaries (*e.g.*, Wright et al. 2012). Transit surveys typically observe binaries, but the question of whether they were searchable to begin with is left for later interpretation. In spectroscopic follow-up of candidate transiting planets, the prevalence of astrophysical false-positives may also lead to a bias against confirmation of transiting planets in binary systems.

Ignoring these complications, in this work we showed that binarity biases transit survey occurrence rates through its effects on star counts and the apparent radii of detected planets. Our results from Sec. 4.2 indicate that binarity could lead to underestimated HJ rates relative to singles by a multiplicative factor of at most ≈ 1.13 (assuming a power-law radius distribution, and that no secondaries host HJs). We later pointed out in Sec. 4.3 that assuming a power-law radius distribution is probably wrong if one wishes to study the HJ rate discrepancy. If we instead assumed a Gaussian radius distribution (following Petigura et al. 2017b), apparent HJ rates are actually *greater* than the true rate around singles. This means that assuming a Gaussian similar to that reported by Petigura et al. (2017b), binarity is unlikely to resolve the HJ rate discrepancy. Of course, such calculations can be at most suggestive – a true resolution of binarity’s effects may require a detailed understanding of the *Kepler* field’s multiplicity statistics, and the mission’s completeness (both for candidate detection and follow-up). For instance, if hot Jupiters are less likely to be confirmed in binary systems, this might bias the rates low.

Does a detected planet orbit the primary or secondary?—Ciardi et al. (2015) studied the effects of stellar multiplicity on the planet radii derived from transit surveys. They modeled the problem for *Kepler* objects of interest by matching a population of binary and tertiary companions to KOI stars, under the assumption that the KIC-

⁴ Petigura et al. (2017b) recently found $\beta = 3.4^{+0.9}_{-0.8}$. This gives an $\approx 25\%$ relative difference, indicating metallicity could account for at least half of the “hot Jupiter rate discrepancy”.

listed stars were the primaries. They then computed planet radius correction factors assuming that *Kepler*-detected planets orbited the primary or companion stars with equal probability (their Sec. 5). Under these assumptions, they found that any given planet’s radius is on average underestimated by a multiplicative factor of ≈ 1.5 . Ziegler et al. (2017) recently reported a similar value for the real population of systems observed by the Robo-AO KOI survey, under the same assumption.

Our models indicate that assuming a detected planet has equal probability of orbiting the primary or secondary leads to an overestimate of binarity’s population-level effects. A planet orbiting the secondary does lead to extreme corrections, but these cases are rare outliers, because the searchable volume for secondaries is so much smaller than that for primaries (see Fig. 1). This means that when high-resolution imaging discovers a binary companion in a system that hosts a detected transiting planet, the planet is much more likely to orbit the primary. This statement is independent of the fact that planets are often confirmed to orbit the primary by inferring the stellar density from the transit duration.

On completeness—**TODO THIS SECTION** This is if some fraction of stars you thought were searchable are not. There are no such stars.

Our expressions for the apparent rate density (Eqs. 2 and 10) do not explicitly mention detection efficiency. If we were to take a smooth detection probability in S/N, we would need to include an explicit detection efficiency term in Eq. 2 to inversely weight for planets that had low probabilities of being detected. However, our “detection probability” is a step function in S/N (Eq. 1). In our model, the maximum searchable distance is only a function of the apparent depth and true system luminosity. Even though the observers are selecting stars based on what they think are “true radii” (bottom row of Fig. 1), the actual searchability of a point-source is determined by apparent radius and system luminosity (Eq. 1; top rows of Fig. 1). Thus at a given apparent radius r_a , the stars (single or binary) that are searchable are exactly the stars that are SEARCHED, since dilution affects the apparent radius and the maximum searchable distance in equal and opposite order.

Utility in future occurrence rate measurements—*TESS* is expected to discover over 10^4 giant planets (Ricker et al. 2014; Sullivan et al. 2015). Though they will be difficult to distinguish from false positives, one possible use of this large sample will to measure an occurrence rate of short-period giant planets with minimal error from counting statistics. Our models suggest that binarity will only become an appreciable fraction of the error budget at the $\approx 10\%$ precision level. This should be sufficient for a rate measurement precise enough to indicate a preference between $\approx 0.5\%$ and $\approx 1\%$ of Sun-like stars hosting hot Jupiters.

6. CONCLUSION

This study presented simple models for the effects of binarity on occurrence rates measured by transit surveys. The simplest of these models (Sec. 2) provided an order-

of-magnitude estimate for how much binarity might affect occurrence rates: by no more than \sim factors of two.

We then gave a general formula for the apparent rate density inferred by an observer ignoring binarity (Eq. 15). As input, this equations requires a volume-limited mass ratio distribution $f(q)$, and also the true rate densities for planets around singles, primaries, and secondaries. The assumptions that make the model tractable are:

1. the transit survey is S/N limited;
2. all the primaries and single stars have identical properties;
3. the properties of the secondaries can vary, but there are known functions $L(M)$ and $R(M)$ that specify a star’s luminosity and radius in terms of its mass;
4. the binarity-ignoring observers assume that binaries have the same properties as the primaries.

Allowing for these conditions, many results follow:

- Assuming a power-law planet radius distribution, and that the same number of planets orbit singles, primaries, and secondaries, binarity influences apparent planet occurrence rates around Sun-like stars at the \sim few percent level from radii $2r_{\oplus} \lesssim r_a \lesssim 17r_{\oplus}$ (Sec. 4.1).
- Assuming a broken-power law planet radius distribution, with [Howard et al. \(2012\)](#)’s exponent at $r > 2r_{\oplus}$ and a constant occurrence at $r < 2r_{\oplus}$, there is a “bump” in the apparent rate density at $r_a < 2r_{\oplus}$, leading to a relative error $\delta\Gamma_0 = |\Gamma_0 - \Gamma_a|/\Gamma_0$ of at most $\approx 50\%$ (Fig. 4). Although this is smaller than current systematic uncertainties on the occurrence rates of Earth-sized planets, this means that binarity could eventually become an important component of the η_{\oplus} error budget.
- Binarity “fills in” gaps in the radius distribution (Fig. 5), by an amount that could affect precise measurements of the depth, width, and slope of [Fulton et al. \(2017\)](#)’s radius gap, in the event that planets are not carefully vetted with high resolution imaging. This will likely become an important consideration for precise measurements of the depth, width, and slope of .
- Binarity does not lead to smaller apparent HJ occurrence rates (Fig. 6). This assumes a Gaussian planet radius distribution, similar to that reported by [Petrigura et al. \(2017b\)](#).

All of these results should be understood as only being strictly applicable when the assumptions listed above are met. Otherwise, they are only suggestive. **TODO: finish**

It was a pleasure discussing this work with F. Dai and T. Barclay.

Software: `numpy` ([Walt et al. 2011](#)), `scipy` ([Jones et al. 2001](#)), `matplotlib` ([Hunter 2007](#)), `pandas` ([McKinney 2010](#))

APPENDIX

A. DERIVATION OF GENERAL FORMULA FOR APPARENT OCCURRENCE RATE

Recall our definition of apparent rate density (Eq. 2). Let us explicitly write $\mathcal{P} = r$ and $\mathcal{S} = M$; R and L are uniquely determined from the assumed stellar mass–radius–luminosity relation. We neglect the dependence on planetary orbital period. The planets with (r_a, M_a) are associated with systems of many different planetary and stellar properties, so $n_{\text{det}}(r_a, M_a)$ is given by the convolution of the true rate density, $\Gamma(r, M)$, and $\mathcal{N}(r_a, M_a; r, M)$, the number (per unit (r_a, M_a)) of searchable stars that give (r_a, M_a) when the true system actually has (r, M) . Mathematically,

$$n_{\text{det}}(r_a, M_a) = \sum_i n_{\text{det}}^i(r_a, M_a) \quad (\text{A1})$$

$$= \sum_i \int dr dM \mathcal{N}_s^i(r_a, M_a; r, M) \cdot \Gamma_i(r, M) \cdot p_{\text{tra}}(r, M), \quad (\text{A2})$$

where i specifies the type of true host stars (0: single, 1: primary, 2: secondary). The problem reduces to the evaluation of

$$\mathcal{N}_s^i(\mathcal{P}_a, \mathcal{S}_a; \mathcal{P}, \mathcal{S}) \quad (\text{A3})$$

for planets around single stars, primaries in binaries, and secondaries in binaries.

Single stars—For $i = 0$,

$$\mathcal{N}_s^0(r_a, M_a; r, M) = \delta(r_a - r) \delta(M_a - M) N_s^0(r, M), \quad (\text{A4})$$

so

$$n_{\text{det}}^0(r_a, M_a) = N_s^0(r_a, M_a) \cdot \Gamma_0(r_a, M_a) \cdot p_{\text{tra}}(r_a, M_a). \quad (\text{A5})$$

Primaries in binaries—The number of primaries with apparent parameters (r_a, M_a) given the true parameters (r, M) is

$$\mathcal{N}_s^1(r_a, M_a; r, M) = \int dq f(q) \mathcal{N}_{s,q}^1(r_a, M_a, q; r, M), \quad (\text{A6})$$

where $f(q)$ is the volume-limited binary mass ratio distribution.

Since we assume $\mathcal{S}_a = \mathcal{S}_1$,

$$\mathcal{N}_{s,q}^1(r_a, M_a, q; r, M) \propto \delta(M_a - M). \quad (\text{A7})$$

In this case, $\mathcal{N}_{s,q}^1$ is non-zero only at $r_a = R_a \sqrt{\delta_{\text{obs}}}$, and the observed depth is

$$\delta_{\text{obs}} = \left[\frac{r}{R(M_a)} \right]^2 \times \frac{L(M_a)}{L_{\text{sys}}(M_a, q)} \equiv \left[\frac{r}{R(M_a)} \right]^2 \times \mathcal{A}(q, M_a)^2 \quad (\text{A8})$$

The normalization of $\mathcal{N}_{s,q}^1$ is given by the number of binaries that are searchable for a signal δ_{obs} and that have the mass ratio q :

$$N_s^0(\delta_{\text{obs}}, L(M_a)) \cdot \frac{n_b}{n_s} \left[\frac{L_{\text{sys}}(M_a, q)}{L(M_a)} \right]^{3/2} = N_s^0(\delta_{\text{obs}}, L(M_a)) \cdot \frac{\text{BF}}{1 - \text{BF}} \cdot \frac{1}{\mathcal{A}(q, M_a)^3}. \quad (\text{A9})$$

Thus,

$$\begin{aligned} \mathcal{N}_{s,q}^1(r_a, M_a, q; r, M) &= N_s^0(\delta_{\text{obs}}, L(M_a)) \cdot \frac{\text{BF}}{1 - \text{BF}} \cdot \frac{1}{\mathcal{A}(q, M_a)^3} \\ &\times \delta[r_a - r\mathcal{A}(q, M_a)] \delta(M_a - M). \end{aligned} \quad (\text{A10})$$

Secondaries in binaries—In this case, $M = qM_1 = qM_a$, so

$$\mathcal{N}_{s,q}^2(r_a, M_a, q; r, M) \propto \delta\left(M_a - \frac{M}{q}\right). \quad (\text{A11})$$

Again \mathcal{N}_s^2 is non-zero only at $r_a = R_a\sqrt{\delta_{\text{obs}}}$, but this time

$$\delta_{\text{obs}} = \left[\frac{r}{R(qM_a)} \right]^2 \times \frac{L(qM_a)}{L_{\text{sys}}(M_a, q)}. \quad (\text{A12})$$

The normalization remains the same as the previous case (we are counting the searchable stars at a given observed depth δ_{obs} , total luminosity of the binary is the same). Thus,

$$\mathcal{N}_s^2(r_a, M_a; r, M) = \int dq f(q) \mathcal{N}_{s,q}^2(r_a, M_a, q; r, M), \quad (\text{A13})$$

where

$$\begin{aligned} \mathcal{N}_s^2(r_a, M_a; r, M; q) &= N_s^0(\delta_{\text{obs}}, L(M_a)) \cdot \frac{\text{BF}}{1 - \text{BF}} \cdot \frac{1}{\mathcal{A}(q, M_a)^3} \\ &\times \delta\left(r_a - r\sqrt{\left[\frac{R(M_a)}{R(qM_a)}\right]^2 \frac{L(qM_a)}{L_{\text{sys}}(M_a, q)}}\right) \delta\left(M_a - \frac{M}{q}\right). \end{aligned} \quad (\text{A14})$$

One might worry in Eq. A14 that we opt to write $\mathcal{N}_s^2 \propto \delta(M_a - M/q)$, rather than $\propto \delta(M_a q - M)$ or some other delta function with the same functional dependence, but a different normalization once integrated. We do this because the delta function in Eq. A14 is defined with respect to the measure dM_a , not dM . This is because \mathcal{N}_s^2 is defined as a number per r_a , per M_a .

Number of detected planets—Marginalizing per Eq. A2, we find

$$\begin{aligned} n_{\text{det}}^0(r_a, M_a) &= \int dr dM \mathcal{N}_s^0(r_a, M_a; r, M) \cdot \Gamma_0(r, M) \cdot p_{\text{tra}}(M) \\ &= N_s^0(\delta_{\text{obs}}, L(M_a)) \cdot \Gamma_0(r_a, M_a) \cdot p_{\text{tra}}(M_a), \end{aligned} \quad (\text{A15})$$

and

$$\begin{aligned} n_{\text{det}}^1(r_a, M_a) &= \int \text{drd}M \mathcal{N}_s^1(r_a, M_a; r, M) \cdot \Gamma_1(r, M) \cdot p_{\text{tra}}(M) \\ &= N_s^0(\delta_{\text{obs}}, L(M_a)) \cdot p_{\text{tra}}(M_a) \cdot \frac{\text{BF}}{1 - \text{BF}} \int \frac{\text{d}q}{\mathcal{A}^4} f(q) \Gamma_1\left(\frac{r_a}{\mathcal{A}}, M_a\right), \end{aligned} \quad (\text{A16})$$

where

$$\mathcal{A}(q, M_a) = \sqrt{\frac{L(M_a)}{L_{\text{sys}}(M_a, q)}}. \quad (\text{A17})$$

Finally,

$$n_{\text{det}}^2(r_a, M_a) = \int \text{drd}M \mathcal{N}_s^2(r_a, M_a; r, M) \cdot \Gamma_2(r, M) \cdot p_{\text{tra}}(M) \quad (\text{A18})$$

$$= N_s^0(\delta_{\text{obs}}, L(M_a)) \cdot \frac{\text{BF}}{1 - \text{BF}} \int \frac{q \text{d}q}{\mathcal{A}^3 \mathcal{B}} f(q) \Gamma_2(r_a/\mathcal{B}, qM_a) p_{\text{tra}}(qM_a), \quad (\text{A19})$$

where

$$\mathcal{B}(q, M_a) = \frac{R(M_a)}{R(qM_a)} \sqrt{\frac{L(qM_a)}{L_{\text{sys}}(M_a, q)}}. \quad (\text{A20})$$

General formula for apparent occurrence rate—Using the results above, the apparent rate density,

$$\Gamma_a(r_a, M_a) = \frac{1}{N_{s,a}(r_a, M_a) p_{\text{tra}}(r_a, M_a)} \times \sum_i n_{\text{det}}^i(r_a, M_a), \quad (\text{A21})$$

evaluates to

$$\begin{aligned} \Gamma_a(r_a, M_a) &= \frac{1}{1 + \mu(\text{BF}, M_a)} \times \left\{ \Gamma_0(r_a, M_a) + \frac{\text{BF}}{1 - \text{BF}} \left[\int \frac{\text{d}q}{\mathcal{A}^4} f(q) \Gamma_1\left(\frac{r_a}{\mathcal{A}}, M_a\right) \right. \right. \\ &\quad \left. \left. + \int \frac{q \text{d}q}{\mathcal{A}^3 \mathcal{B}} f(q) \Gamma_2\left(\frac{r_a}{\mathcal{B}}, qM_a\right) \frac{R(qM_a)}{R(M_a)} q^{-1/3} \right] \right\}. \end{aligned} \quad (\text{A22})$$

This equation is used to derive Eqs. 23 and 27, and is numerically integrated to create Figs. 4, 5, and 6. It is validated in limits in which it is possible to write down the answer (e.g., Eq. 10), and also against a Monte Carlo realization of the twin binary models (Secs. 2.1 and 4.1).

Table 1. Occurrence rates of hot Jupiters (HJs) about FGK dwarfs, as measured by radial velocity and transit surveys.

Reference	HJs per thousand stars	HJ Definition
Marcy et al. (2005)	12±2	$a < 0.1$ AU; $P \lesssim 10$ day
Cumming et al. (2008)	15±6	—
Mayor et al. (2011)	8.9±3.6	—
Wright et al. (2012)	12.0±3.8	—
Gould et al. (2006)	3.1 ^{+4.3} _{-1.8}	$P < 5$ day
Bayliss & Sackett (2011)	10 ⁺²⁷ ₋₈	$P < 10$ day
Howard et al. (2012)	4±1	$P < 10$ day; $r_p = 8 - 32r_\oplus$; solar subset ^a
—	5±1	solar subset extended to $Kp < 16$
—	7.6±1.3	solar subset extended to $r_p > 5.6r_\oplus$.
Moutou et al. (2013)	10±3	<i>CoRoT</i> average; $P \lesssim 10$ day, $r_p > 4r_\oplus$
Petigura et al. (2018, in prep)	5.7 ^{+1.4} _{-1.2}	$r_p = 8 - 24r_\oplus$; $P = 1 - 10$ day; CKS stars ^b
Santerne et al. (2018, in prep)	9.5±2.6	<i>CoRoT</i> galactic center
—	11.2±3.1	<i>CoRoT</i> anti-center

NOTE— The first four studies use data from radial velocity surveys; the rest are based on transit surveys. Many of these surveys selected different stellar samples. “—” denotes “same as above”.

^a Howard et al. (2012)’s “solar subset” was defined as *Kepler*-observed stars with $4100 \text{ K} < T_{\text{eff}} < 6100 \text{ K}$, $Kp < 15$, $4.0 < \log g < 4.9$. They required signal to noise > 10 for planet detection.

^b Petigura et al. (2018, in prep)’s planet sample includes all KOIs with $Kp < 14.2$, with a statistically insignificant number of fainter stars with HZ planets and multiple transiting planets. Their stellar sample begins with Mathur et al. (2017)’s catalog of 199991 *Kepler*-observed stars. Successive cuts are: $Kp < 14.2$ mag, $T_{\text{eff}} = 4700 - 6500 \text{ K}$, and $\log g = 3.9 - 5.0$ dex, leaving 33020 stars.

REFERENCES

- Adams, E. R., Ciardi, D. R., Dupree, A. K., et al. 2012, *The Astronomical Journal*, 144, 42
- Adams, E. R., Dupree, A. K., Kulesa, C., & McCarthy, D. 2013, *The Astronomical Journal*, 146, 9
- Bakos, G. ., Csabry, Z., Penev, K., et al. 2013, *PASP*, 125, 154
- Baranec, C., Ziegler, C., Law, N. M., et al. 2016, *The Astronomical Journal*, 152, 18
- Batalha, N. M., Borucki, W. J., Koch, D. G., et al. 2010, *The Astrophysical Journal Letters*, 713, L109
- Bayliss, D. D. R., & Sackett, P. D. 2011, *The Astrophysical Journal*, 743, 103
- Burke, C. J., Christiansen, J. L., Mullally, F., et al. 2015, *The Astrophysical Journal*, 809, 8
- Cartier, K. M. S., Gilliland, R. L., Wright, J. T., & Ciardi, D. R. 2015, *The Astrophysical Journal*, 804, 97
- Ciardi, D. R., Beichman, C. A., Horch, E. P., & Howell, S. B. 2015, *The Astrophysical Journal*, 805, 16
- Cumming, A., Butler, R. P., Marcy, G. W., et al. 2008, *Publications of the Astronomical Society of the Pacific*, 120, 531
- Dong, S., & Zhu, Z. 2013, *The Astrophysical Journal*, 778, 53

- Dong, S., Zheng, Z., Zhu, Z., et al. 2014, *The Astrophysical Journal Letters*, 789, L3
- Dressing, C. D., Adams, E. R., Dupree, A. K., Kulesa, C., & McCarthy, D. 2014, *The Astronomical Journal*, 148, 78
- Dressing, C. D., & Charbonneau, D. 2015, *ApJ*, 807, 45
- Everett, M. E., Barclay, T., Ciardi, D. R., et al. 2015, *The Astronomical Journal*, 149, 55
- Fischer, D. A., & Valenti, J. 2005, *The Astrophysical Journal*, 622, 1102
- Foreman-Mackey, D., Hogg, D. W., & Morton, T. D. 2014, *The Astrophysical Journal*, 795, 64
- Fressin, F., Torres, G., Charbonneau, D., et al. 2013, *The Astrophysical Journal*, 766, 81
- Fulton, B. J., Petigura, E. A., Howard, A. W., et al. 2017, *The Astronomical Journal*, 154, 109
- Furlan, E., & Howell, S. B. 2017, *arXiv:1707.01942 [astro-ph]*
- Furlan, E., Ciardi, D. R., Everett, M. E., et al. 2017, *The Astronomical Journal*, 153, 71
- Gilliland, R. L., Cartier, K. M. S., Adams, E. R., et al. 2015, *The Astronomical Journal*, 149, 24
- Gould, A., Dorsher, S., Gaudi, B. S., & Udalski, A. 2006, *Acta Astronomica*, 56, 1
- Günther, M. N., Queloz, D., Demory, B.-O., & Bouchy, F. 2017, *Monthly Notices of the Royal Astronomical Society*, 465, 3379
- Guo, X., Johnson, J. A., Mann, A. W., et al. 2017, *The Astrophysical Journal*, 838, 25
- Holman, M. J., & Wiegert, P. A. 1999, *The Astronomical Journal*, 117, 621
- Horch, E. P., Howell, S. B., Everett, M. E., & Ciardi, D. R. 2012, *The Astronomical Journal*, 144, 165
- . 2014, *The Astrophysical Journal*, 795, 60
- Howard, A. W., Marcy, G. W., Bryson, S. T., et al. 2012, *The Astrophysical Journal Supplement Series*, 201, 15
- Howell, S. B., Everett, M. E., Sherry, W., Horch, E., & Ciardi, D. R. 2011, *The Astronomical Journal*, 142, 19
- Hunter, J. D. 2007, *Computing in Science & Engineering*, 9, 90
- Johnson, J. A., Petigura, E. A., Fulton, B. J., et al. 2017, *arXiv:1703.10402 [astro-ph]*
- Jones, E., Oliphant, T., Peterson, P., et al. 2001, *Open source scientific tools for Python*
- Kraus, A. L., Ireland, M. J., Huber, D., Mann, A. W., & Dupuy, T. J. 2016, *The Astronomical Journal*, 152, 8
- Law, N. M., Morton, T., Baranec, C., et al. 2014, *The Astrophysical Journal*, 791, 35
- Lillo-Box, J., Barrado, D., & Bouy, H. 2012, *Astronomy and Astrophysics*, 546, A10
- . 2014, *Astronomy and Astrophysics*, 566, A103
- Marcy, G., Butler, R. P., Fischer, D., et al. 2005, *Progress of Theoretical Physics Supplement*, 158, 24
- Mathur, S., Huber, D., Batalha, N. M., et al. 2017, *The Astrophysical Journal Supplement Series*, 229, 30
- Mayor, M., Marmier, M., Lovis, C., et al. 2011, *ArXiv e-prints*, 1109, *arXiv:1109.2497*
- McKinney, W. 2010, in *Proceedings of the 9th Python in Science Conference*, ed. S. van der Walt & J. Millman, 51
- Moutou, C., Deleuil, M., Guillot, T., et al. 2013, *Icarus*, 226, 1625
- Owen, J. E., & Wu, Y. 2017, *arXiv:1705.10810 [astro-ph]*
- Pepper, J., & Gaudi, B. S. 2005, *The Astrophysical Journal*, 631, 581
- Pepper, J., Gould, A., & Depoy, D. L. 2003, *Acta Astronomica*, 53, 213
- Petigura, E. A., Howard, A. W., & Marcy, G. W. 2013, *Proceedings of the National Academy of Science*, 110, 19273

- Petigura, E. A., Howard, A. W., Marcy, G. W., et al. 2017a, [arXiv:1703.10400 \[astro-ph\]](#)
- Petigura, E. A., Marcy, G. W., Winn, J. N., et al. 2017b, [arXiv:1712.04042 \[astro-ph\]](#)
- Raghavan, D., McAlister, H. A., Henry, T. J., et al. 2010, *The Astrophysical Journal Supplement Series*, 190, 1
- Ricker, G. R., Winn, J. N., Vanderspek, R., et al. 2014, *Journal of Astronomical Telescopes, Instruments, and Systems*, 1, 014003
- Santos, N. C., Israelian, G., & Mayor, M. 2004, *Astronomy and Astrophysics*, 415, 1153
- Sullivan, P. W., Winn, J. N., Berta-Thompson, Z. K., et al. 2015, *The Astrophysical Journal*, 809, 77
- Van Eylen, V., Agentoft, C., Lundkvist, M. S., et al. 2017, [arXiv:1710.05398 \[astro-ph\]](#)
- Walt, S. v. d., Colbert, S. C., & Varoquaux, G. 2011, *Computing in Science & Engineering*, 13, 22
- Wang, J., Fischer, D. A., Horch, E. P., & Huang, X. 2015a, *The Astrophysical Journal*, 799, 229
- Wang, J., Fischer, D. A., Horch, E. P., & Xie, J.-W. 2015b, *The Astrophysical Journal*, 806, 248
- Wang, J., Fischer, D. A., Xie, J.-W., & Ciardi, D. R. 2015c, *The Astrophysical Journal*, 813, 130
- Wang, J., Xie, J.-W., Barclay, T., & Fischer, D. A. 2014, *The Astrophysical Journal*, 783, 4
- Wright, J. T., Marcy, G. W., Howard, A. W., et al. 2012, *The Astrophysical Journal*, 753, 160
- Youdin, A. N. 2011, *ApJ*, 742, 38
- Ziegler, C., Law, N. M., Baranec, C., et al. 2017, [arXiv:1712.04454 \[astro-ph\]](#)