

PERIODIC VARIABILITY FROM THE 2MASS CALIBRATION SCANS

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ABSTRACT

data comes from 2mass-calpswdb, have over 100 million photometric measurements, totaling light curves for 113,030 point objects. we have run supersmooth period folding, ala MACHO, on every light curve. By hand verification used to check. many are found periodic, XX binaries, YY radial pulsators. These include the best sampled light curves for an RR Lyr and Cepheid variables in the NIR, and provide an important benchmark for modeling of these stars. An additional ZZ quasi-periodic variable stars are found. we have also characterized the variability properties as functions of color, finding nearly all objects have “grey” variability, where the amplitude is independent of wavelength chosen.

Subject headings: variability, surveys, NIR

1. INTRODUCTION

Arguably the two greatest advancements in optical and infrared photometric surveys within the past two decades have been 1) the advent of 1% accurate multi-wavelength photometry for contiguous portions of the sky exceeding $\Omega \sim 10^5$ deg², and 2) time resolved surveys spanning days to years baselines. As both the temporal and spatial domains are explored, new astrophysical phenomena are uncovered, while the previously rare events are placed in a statistical context. Naturally, the next generation survey will exploit both of these domains simultaneously, producing deep multi-wavelength surveys with both unprecedented spatial and temporal coverage. In order to prepare for these future surveys, which will contain orders of magnitude larger numbers of variable sources, we must hone our skills using existing catalogs/databases. This will precipitate the development of new techniques, and inform the targeting, cadence, and observing strategies of the many future time domain surveys.

The Two Micron All Sky Survey (2MASS; Skrutskie et al. 2006) imaged the full sky in three, simultaneously obtained, near-infrared (NIR) bands. Survey operations were conducted between 1997 and 2001, using a northern and southern telescope. Photometric calibration for 2MASS was accomplished using repeated observations of 35 selected fields, which were spaced across the sky (Nikolaev et al. 2000). An additional 5 tiles were imaged around the Large and Small Magellanic Clouds during the last year of operation. Each calibration field covered an area of approximately 8'.5 (RA) \times 1° (Dec). The 35 standard calibration fields were scanned ranging between 562 and 3,692 times over the 4 year period, yielding some of the best NIR light curves yet produced. The placement of these tiles is shown in Figure 1, and they contain light curves for approximately 110,000 point source objects.

This unique dataset has been used for a handful of studies to date. Plavchan et al. (2008b) characterized many of the details in analyzing this unique dataset for

time domain studies. They also mined these data for periodic objects, finding 3 new M dwarf eclipsing binaries. Plavchan et al. (2008a) studied the 131 day periodic object, 2MASS J16271848–2429059, revealing a possible three-body YSO system. Sarajedini et al. (2009) utilized the images from multiple scans of calibration tile 90067 to produce a NIR color-magnitude diagram for the open cluster M67 that probed $\gtrsim 3$ magnitudes deeper than the standard 2MASS point source catalog. Becker et al. (2008) discovered and characterized a 2.6 d M dwarf binary, 2MASS J01542930+0053266. This system was located in both the 2MASS calibration tile 90004, and the “Stripe 82” region of the Sloan Digital Sky Survey (SDSS; York et al. 2000) footprint, which yielded an 8-band light curve ($ugrizJHK_s$) for the binary. Using 16 of the calibration tiles that overlapped the SDSS footprint, Davenport et al. (2012) produced some of the first constraints on the properties of M dwarf flares in red optical and NIR bandpasses.

MENTION THESE TWO? (P. Plavchan 2012 submitted) an analysis of a 92 day periodic YSO. (J. Parks 2012 submitted) a detailed study of the ρ Ophiuchus star forming region tile.

ADD PARAGRAPH PROMPTING USE OF 2MASS AS CENSUS FOR VARIABLES – TAKE IDEAS FROM PROPOSAL

In this paper we present a census of the periodic variable objects from this 4 year NIR survey. The data and period-finding methodology is outlined in §2. We describe the selection and classification of binary stars and radial pulsating objects in §3. Quasi-periodic and other large amplitude variables that were recovered are discussed in §4. We examine the general variability characteristics of point sources in the 2MASS calibration scans in §5, and concluding remarks are given in §6.

2. DATA

lots of testing, characterization of the Cal-PSWDB done in (Plavchan 2006; Plavchan et al. 2008b) Each night during 2MASS operations, the survey telescopes were directed at one of the calibration fields once per hour (before 1997 October 11 UT two fields were observed every 2 hours). During each visit, six consecutive

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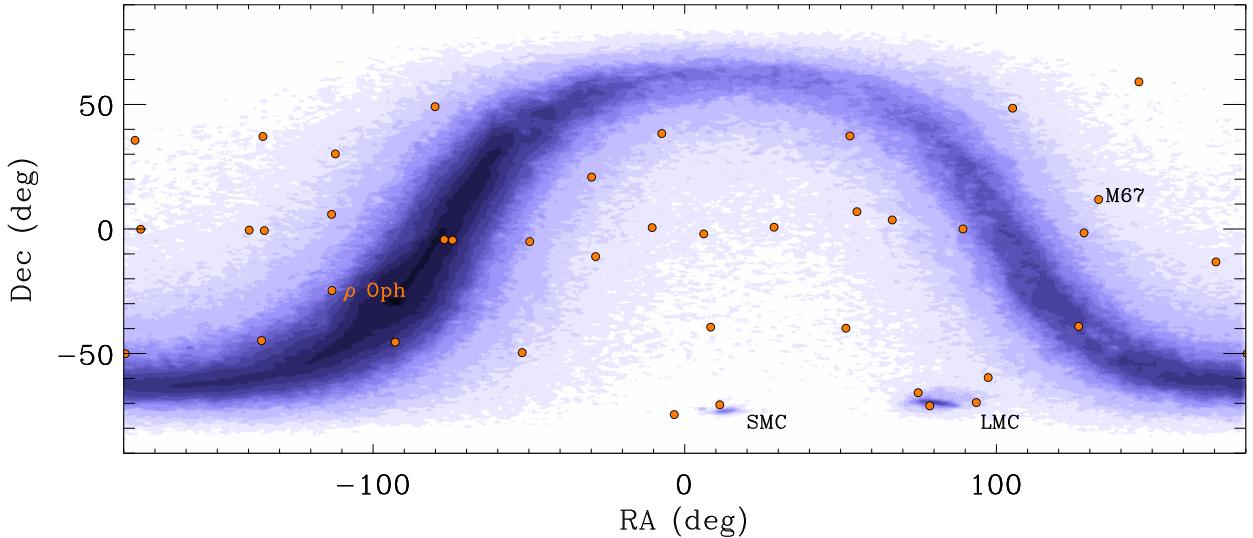


FIG. 1.— Spatial distribution of the 40 Cal-PSWDB tiles (orange circles), with titles for a few notable tiles labeled. Background contours denote increasing density from a sample of 3 million randomly drawn point sources from the 2MASS point source catalog, and trace the Galactic plane, bulge, and the Large and Small Magellanic Clouds.

scans of the field were made in alternating declination directions in ~ 10 minutes of elapsed real time (a scan group). Each scan in the set of six was offset from the preceding scan by $\sim 5''$ in R.A. to avoid systematic pixel effects. The calibration fields were observed using the same freeze-frame scanning strategy used for the main survey that yielded a net 7.8 s exposure on the sky per scan. Over the course of the 2MASS survey, between 562 and 3692 independent observations were made of each of the 35 calibration fields. Table 1 presents a list of these fields and their aggregate properties.

The raw imaging data from each scan of a 2MASS calibration field were reduced using the same automated data processing system used to process the survey observation data (Cutri 2006) (Cutri et al. 2006, Section IV). The reduction process detected and extracted source positions and photometry for all objects in the images from each scan. Measurements of the standard stars in each field were used to determine the nightly photometric zero-point solutions as a function of time, and seasonal atmospheric coefficients. All source extractions from all scans were loaded into the 2MASS Calibration Point Source Working Database (Cal-PSWDB). This database contains over 191 million source extractions derived from 73,230 scans of the 35 calibration fields. Further descriptions of the Cal-PSWDB and its properties can be found in Cutri et al. (2006).

we ran the data through optics, got XX objects.
each light curve was processed with supersmsoother
objects with periods in at least 2 bands that were not
obviously aliases were grabbed (about 2000 objects)
each light curve was examined by eye, about 250 were
bona fide good periodic objects

3. PERIODIC VARIABLES

things that SUPERSMOOTH said were periodic in at least 2 bands we looked at, returned 2100 interesting objects. we searched all 2000 light curves of periodic things by eye to pick out the real ones. quite a few

are noise. found correct periods by hand, trying folding objects at 1/2, 2, 2/3 etc. also folded at all periods determined for each band.

we also put the objects in to 4 classification bins by hand: obvious binaries with distinct eclipses, objects with very RR Lyr looking saw-tooth light curves, and sinusoidals which could be either W UMa or RR Lyr (a/b?). we also noted quasi-periodic or long period objects whose periods were not well sampled enough to accurately determine. those could be things like Miras, etc. this recovered 160 binaries, 8 pulsators, 25 other things.

to quantitatively separate these things, need a more impartial method. Eclipsing binaries and pulsating variables can be separated and classified by fitting Fourier modes to the light curves (Pojmanski 2002; Nefs et al. 2012).. we used similar methods as other authors, using the IDL FOURFIT package by Buie (online CITE/footnote).

the binary type separation has been well-explored, and we use previously established equations (CITE Becker? for example). this puts them in 3 categories, which had excellent agreement with our initial by-eye analysis. we recover XX detached, YY semi-detached objects. contact binaries are called W UMa type, have very sinusoidal light curves. these are difficult to distinguish from rr lyr. sometimes they have color variations, but sometimes not if the mass ratio is near 1.

the separation between binary stars and rr lyrae is more difficult. some types of RR Lyr look very similar to binaries, most especially RR Lyr type a (CHECK ON THIS). to improve this, we need other types of Fourier cuts to quantify the asymmetry in the light curves. there is some previous work on this ((?)) but we found their cuts didn't make sense, possibly due to different fitting codes (?) or different wavelength(?).

using our by-eye categorization for objects, 8 for sure pulsators, we explored many combinations of Fourier parameter space. a4/a2 had been suggested previously as

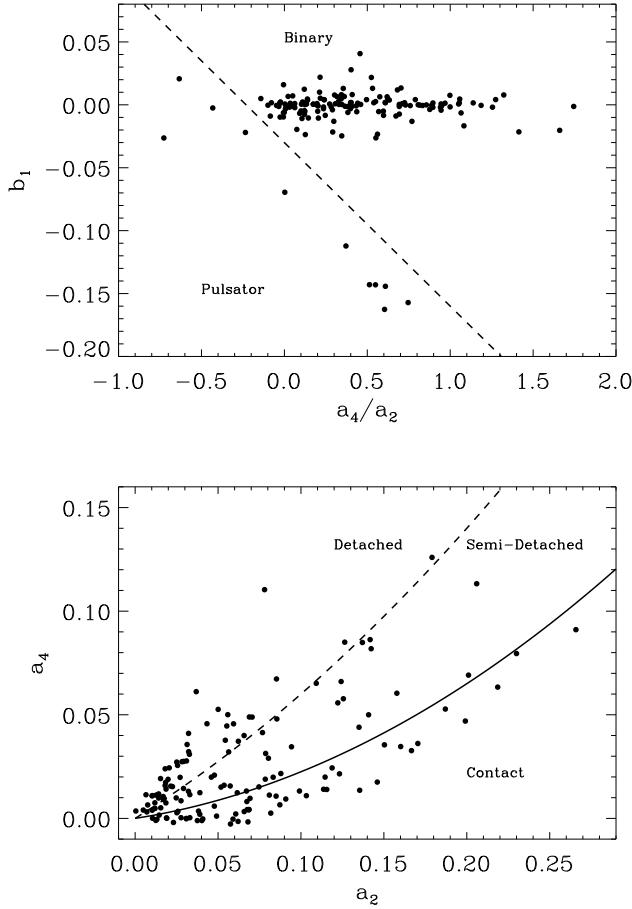


FIG. 2.— Fourier modes used to select between eclipsing binaries and pulsators (top) and different classes of eclipsing binaries (bottom).

a good space to search in (Pojmanski 1997). b_1 makes sense, it is the first component of cosine asymmetry from the sinusoidal shape. binaries with eccentricity near 0 should not have any strong b modes.

To separate binaries from pulsators, we cut on:

$$b_1(\text{binary}) \geq -0.13 \left(\frac{a_4}{a_2} \right) - 0.3 \quad (1)$$

and different configurations of binaries can be determined using:

$$a_4(\text{detached}) \geq a_2(0.5 + a_2) \quad (2)$$

$$a_4(\text{contact}) \leq a_2(0.125 + a_2) \quad (3)$$

4. BINARIES

We found 100 binaries, need a histogram of their periods. We recover nearly all the binaries from Plavchan's early work (8/23 missed so far)

match to new paper by J Parks (2012)

Using our by-hand classification, we can define the best line of delineation between contact and detached eclipsing binaries in this Fourier space. Rucinski (1997) used $a_4 < a_2(0.3 + a_2)$ to separate these two populations using OGLE light curves. We had 64 detached, 72 detached by hand classified. We used a brute-force approach to solve

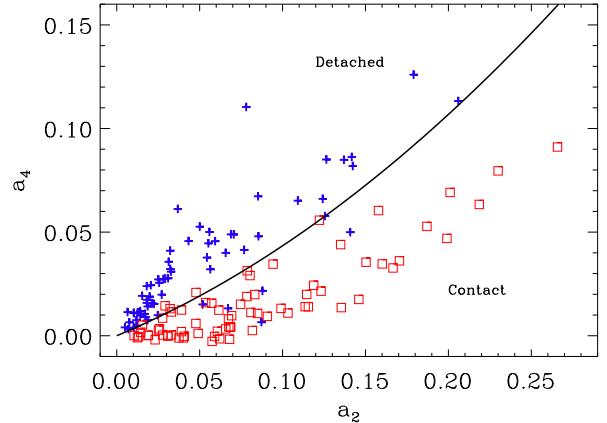


FIG. 3.— My classification of different classes of eclipsing binaries (bottom). separation given in text

for the coefficient C in

$$a_4 = a_2(C + a_2), \quad (4)$$

calculating the fraction of each population correctly classified at each iteration. We used values of c ranging from 0 to 1, in steps of 0.001. The best degree of separation was found for $C = 0.334$, yielding 95% completeness for contact binaries, and 92% completeness for detached systems.

the final sample of classified objects were: XX contact binaries, YY semi-detached, ZZ detached, 12 pulsators, and 25 quasi-periodic oddballs. a selection of the LPV/QPVs are given in Figure 16 (MOVE UP). this matched with total rates of these objects from other surveys? Check out OGLE or ASAS, have they got similar proportions of binaries and rr Lyr?

for every binary, estimate the spectral type based on Davenport (2013 in prep) color-locus. use SDSS colors for stars w/ matches, and WISE colors for everything else (hopefully). this provides a much improved ability to select dM's, rather than a broad-brush color box previously used. We find XX eclipsing binaries with M dwarf colors ($g-i$ between 2 and 3, or whatever). are any of these detached binaries? if so, gold mine. 5 of these are new, we have missed 5 from Plavchan work. Since very in depth effort been put towards rho Oph field, havent focused on these fields.

55 contact, 43 semi-detached, 48 detached binaries based on the equations

examples of each type

5. PULSATORS

Infrared observations of radial pulsating stars was first performed by Wisniewski & Johnson (1968), and detailed characterization was done for Cepheid variables by McGonegal et al. (1982) and for RR Lyr variables by Longmore et al. (1985)

K -band templates first derived in Jones et al. (1996)

RR Lyr type variables belong broadly to two classes: RRab and RRc, defined by their light curve morphologies (Bailey 1902)

Sollima et al. (2008) produced JHK_s light curves for the prototypical star, RR Lyr.

limited numbers found in the NIR, usually many few

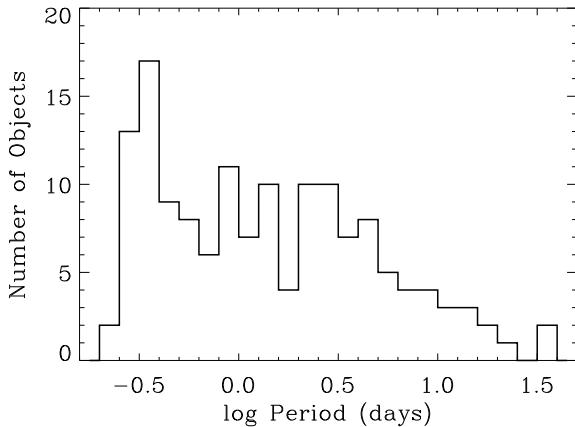


FIG. 4.— Distribution of periods for objects selected as binaries. Note the sharp cutoff at $\log P \sim -0.6$, $P \sim 0.25$ days.

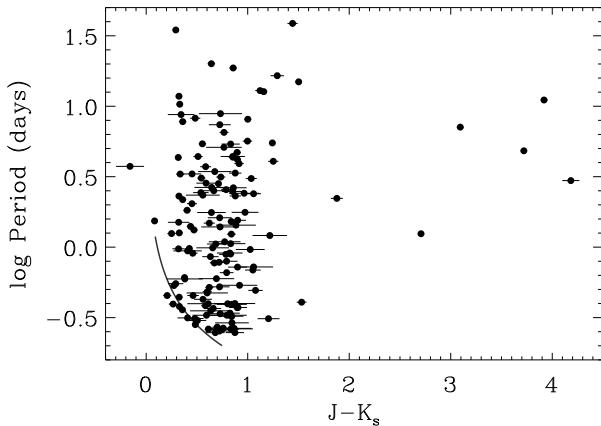


FIG. 5.— Period versus median colors for objects selected as binaries. The photometry was not corrected for reddening. The power law short period contact binary limit from Deb & Singh (2011) is shown for comparison (solid grey line). NEW COLOR FOR LINE, MAKE SYMBOLS DIFFERENT USING CLASSIFICATION FROM FIGURE 2

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numbers of epochs and phase coverage (e.g. Del Principe et al. 2005)

— outline of section — these NIR light curves provide benchmarks for several classes of radial pulsating periodic variables. phase coverage is excellent for most of these objects

we have the best cepheid NIR light curve, best regular RR Lyr, a good δ Scu, and probably some RR a and RR b types (need to look up difference). this list was matched up against SIMBAD, only XX were previously known about.

No Blazko variables - though should check the known sample to see if any have it and don't show it here.

the list is given in Table X, including distances computed using de-reddened and period-mag relations (CITE)

5.1. Effects of Shocks

the light curves can also be folded in to the color-mag space, and medians show kinks. we do this for all the objects.

6. NON-PERIODIC VARIABLES

Plavchan and student have done lots of work on this for ρ Oph field. We will only add other things that were identified by hand

7. CONCLUSIONS

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2MASS Variability

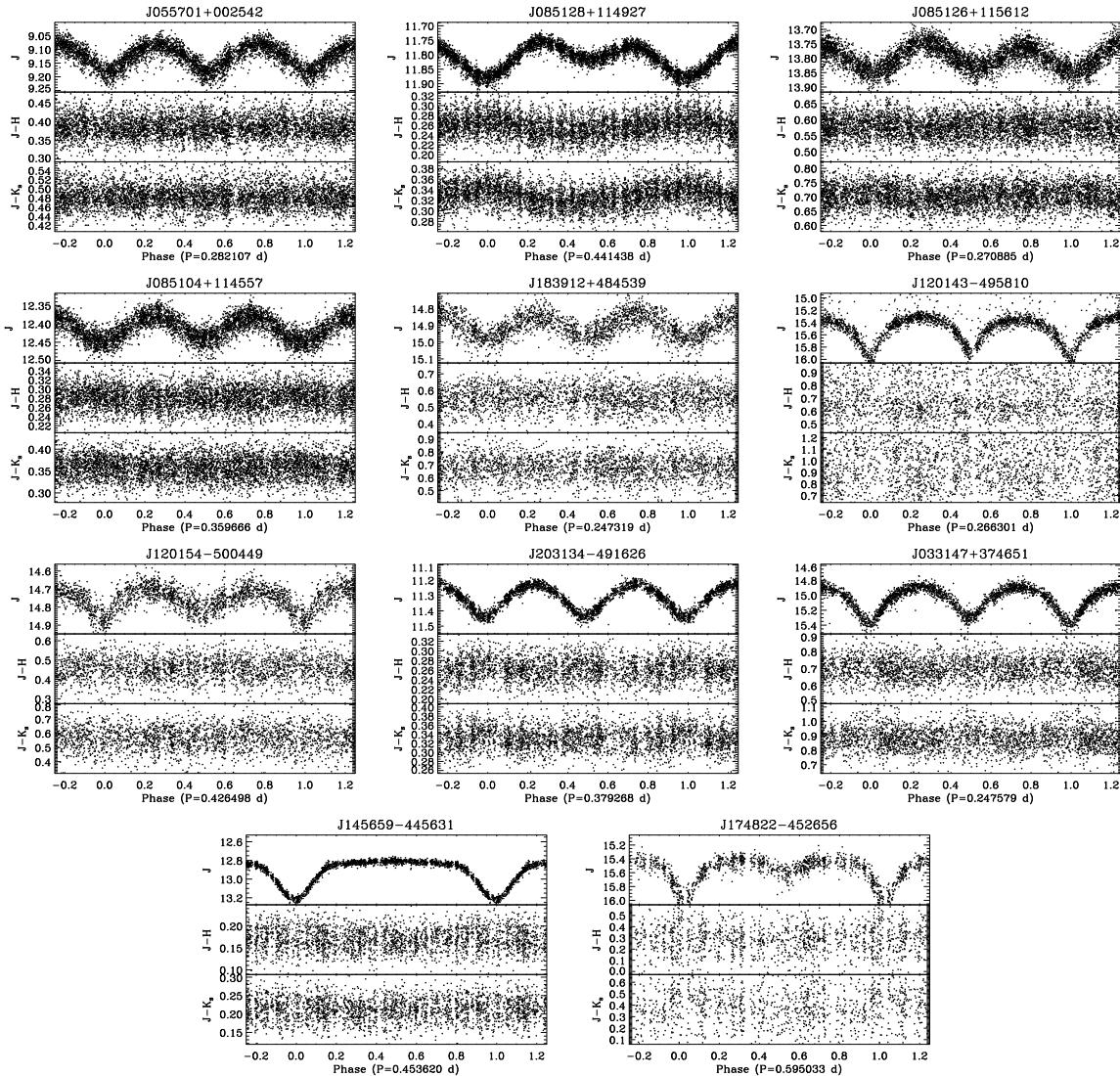


FIG. 6.— contact binaries

TABLE 1

THE CATALOG OF QUASI-PERIODIC VARIABLES. (THIS TEX FILE NEEDS TO BE REVISED TO ONLY INCLUDE THE ACTUALLY INTERESTING OBJECTS, MOST ARE ALREADY SHOWN)

ObjectID (hhmmss+ddmmss)	FieldID	RA (deg)	Dec (deg)	$\langle J \rangle$ (mag)	$\langle H \rangle$ (mag)	$\langle K_s \rangle$ (mag)	# epochs	σ_J (mag)	σ_H (mag)	σ_K (mag)
J162659-243556	90009	246.74608	-24.59911	16.40	13.45	11.86	1543	0.01401	0.00569	0.00182
J162722-244807	90009	246.84578	-24.80195	10.92	9.83	9.34	1581	-0.00000	-0.00000	0.00046
J162718-245453	90009	246.82655	-24.91494	11.42	10.54	9.95	1567	0.07128	0.04564	0.02912

TABLE 2

THE CATALOG OF BINARY STARS, SELECTED USING FOURIER MODES. THE FULL VERSION OF THIS CATALOG (146 ENTRIES) WILL BE AVAILABLE ONLINE.

ObjectID (hhmmss+ddmmss)	FieldID	RA (deg)	Dec (deg)	Period (days)	$\langle J \rangle$ (mag)	$\langle H \rangle$ (mag)	$\langle K_s \rangle$ (mag)	# epochs	a_2	a_4	b_1
J015452+011053	90004	28.72067	1.18153	0.37205	13.62	12.94	12.72	2956	-0.00098	0.00166	0.01248
J015429+005327	90004	28.62207	0.89091	2.63902	15.51	14.84	14.65	2969	0.01922	-0.00187	0.01532
J162709-243408	90009	246.78799	-24.56892	4.83045	12.63	10.25	8.91	1580	0.03721	-0.00837	0.06224

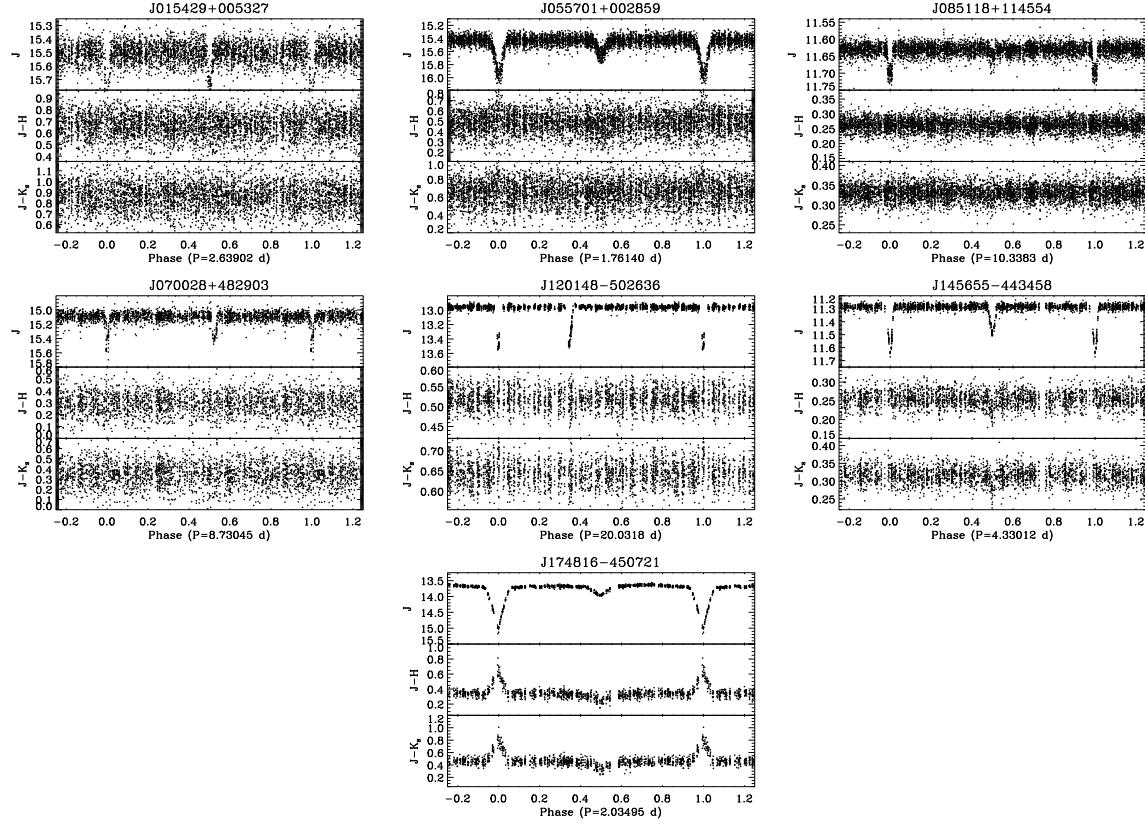
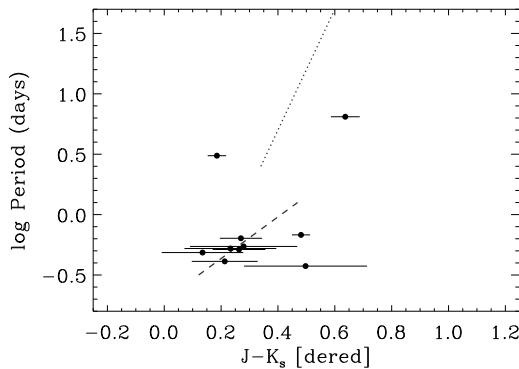


FIG. 7.— detached binaries

FIG. 8.— 17 radial pulsator variables. dashed line is theoretical prediction from Catelan et al. (2004) with $Z=0.001$ (no discernible dependence on Z) dotted line is for cepheids, from ...

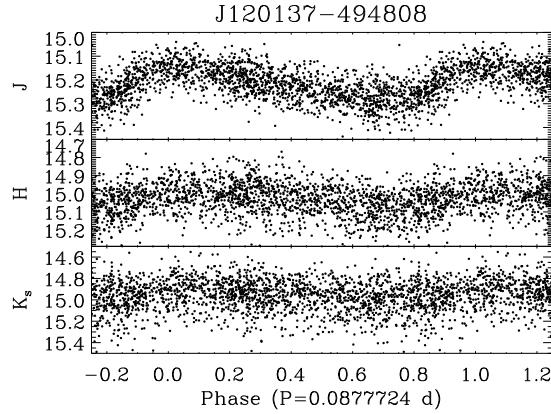
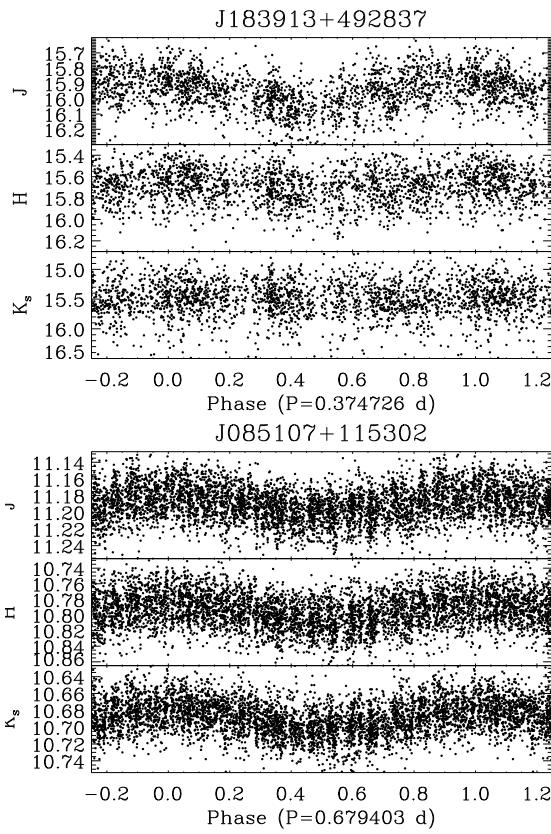
FIG. 9.— δ Scu variable

FIG. 10.— Pulsators with RRc-type light curves.

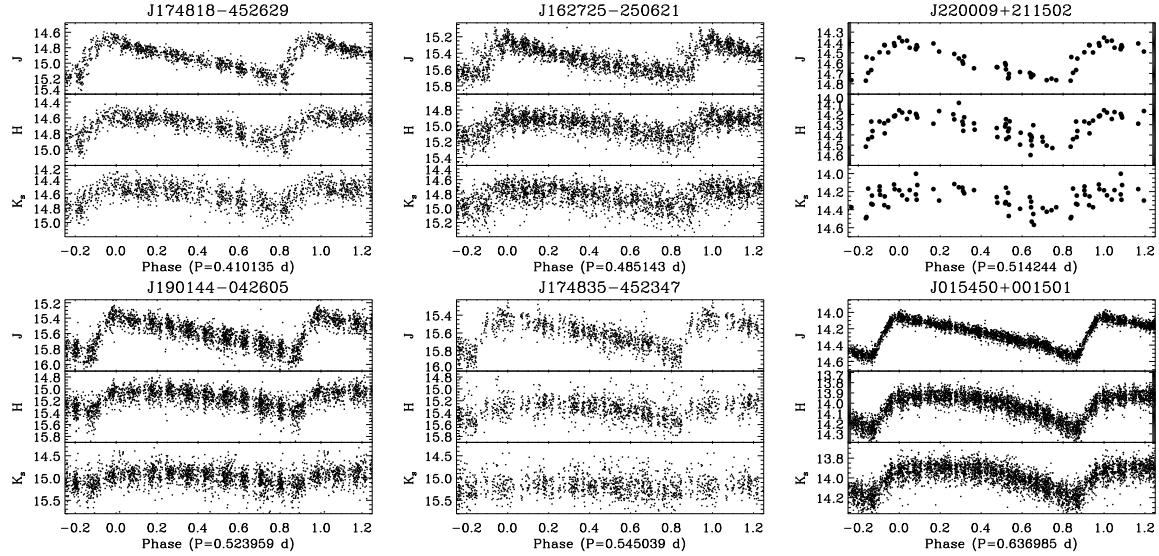


FIG. 11.— RRab-type variables

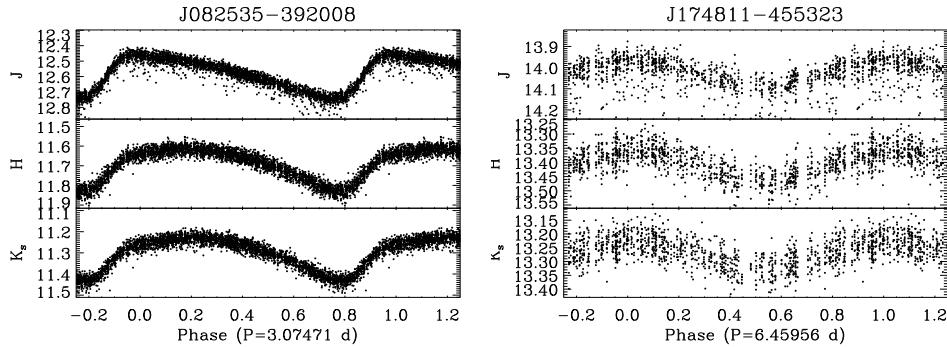


FIG. 12.— Pulsators (cepheids) with periods longer than 2 days

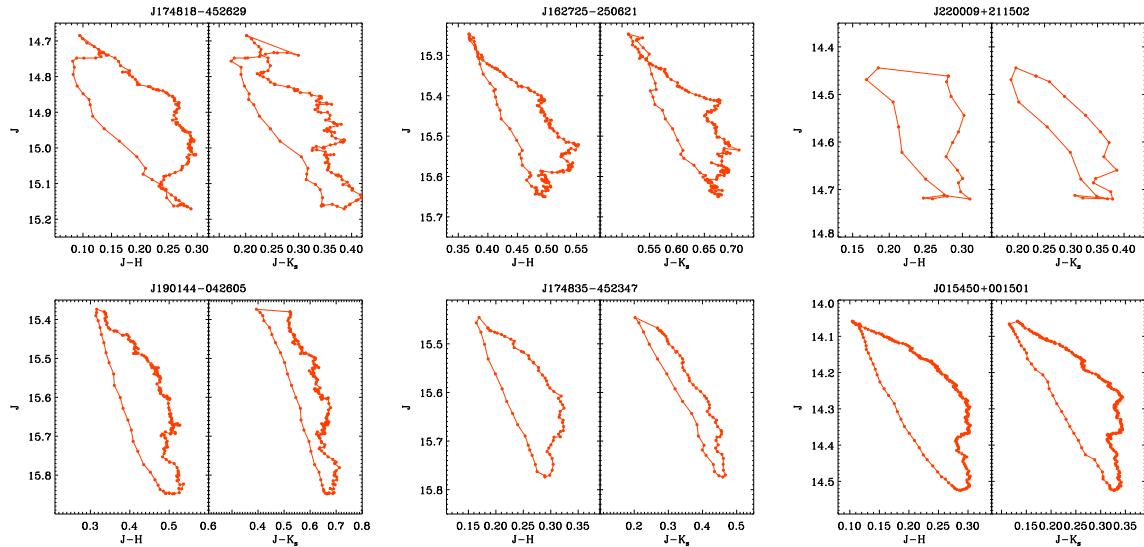


FIG. 13.— CMD of RRab pulsators

2MASS Variability

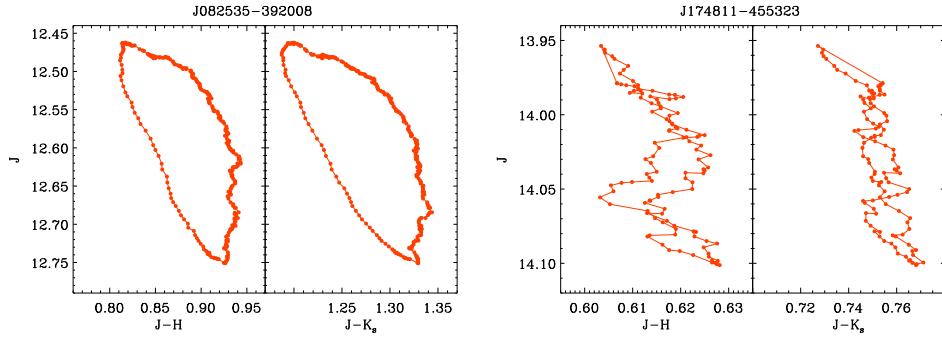


FIG. 14.— CMD of Pulsators (cepheids) with periods longer than 2 days

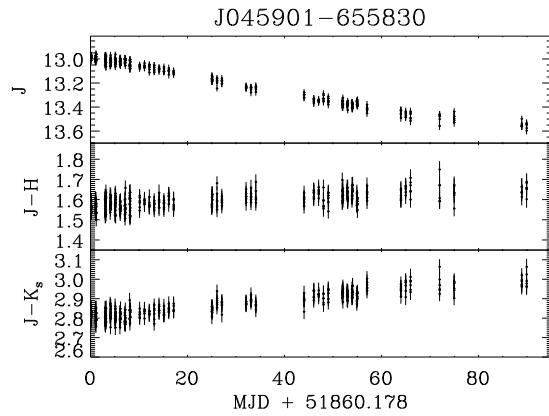


FIG. 15.— LMC, 377 epochs over 90 days, looks kinda like nova of some sort, no other known observations (that i've found)

TABLE 3
THE CATALOG OF RADIAL PULSATING TYPE PERIODIC VARIABLES, SELECTED USING FOURIER MODES.

ObjectID (hhmmss+ddmmss)	FieldID	RA (deg)	Dec (deg)	Period (days)	$\langle J \rangle$ (mag)	$\langle H \rangle$ (mag)	$\langle K_s \rangle$ (mag)	# epochs	a_2	a_4	b_1
J120137-494808	90217	180.40575	-49.80240	0.08777	15.22	15.03	14.95	1685	0.03441	-0.14304	0.06262
J183913+492837	90182	279.80688	49.47710	0.37473	15.96	15.68	15.43	1682	0.01161	-0.14297	0.02263
J174818-452629	90279	267.07571	-45.44157	0.41014	14.93	14.68	14.60	977	0.00012	-0.00248	-0.00028
J162725-250621	90009	246.85564	-25.10589	0.48514	15.46	14.98	14.82	1579	-0.00560	0.02069	0.00883
J220009+211502	92409	330.04028	21.25065	0.51424	14.60	14.35	14.29	41	-0.00234	-0.02196	0.00994
J190144-042605	90808	285.43744	-4.43475	0.52396	15.61	15.14	14.97	1879	-0.00112	-0.02633	0.00154
J174835-452347	90279	267.14786	-45.39640	0.54504	15.63	15.35	15.22	972	0.02224	-0.16258	0.03681
J015450+001501	90004	28.70900	0.25038	0.63698	14.27	14.01	13.98	2972	0.02093	-0.11219	0.05638
J085107+115302	90067	132.78021	11.88394	0.67940	11.19	10.80	10.69	3692	0.03375	-0.15717	0.04523
J082535-392008	90312	126.39848	-39.33574	3.07471	12.57	11.66	11.28	3501	0.00012	-0.06949	0.04578
J174811-455323	90279	267.04904	-45.88993	6.45956	14.02	13.40	13.27	968	0.03683	-0.14425	0.06046

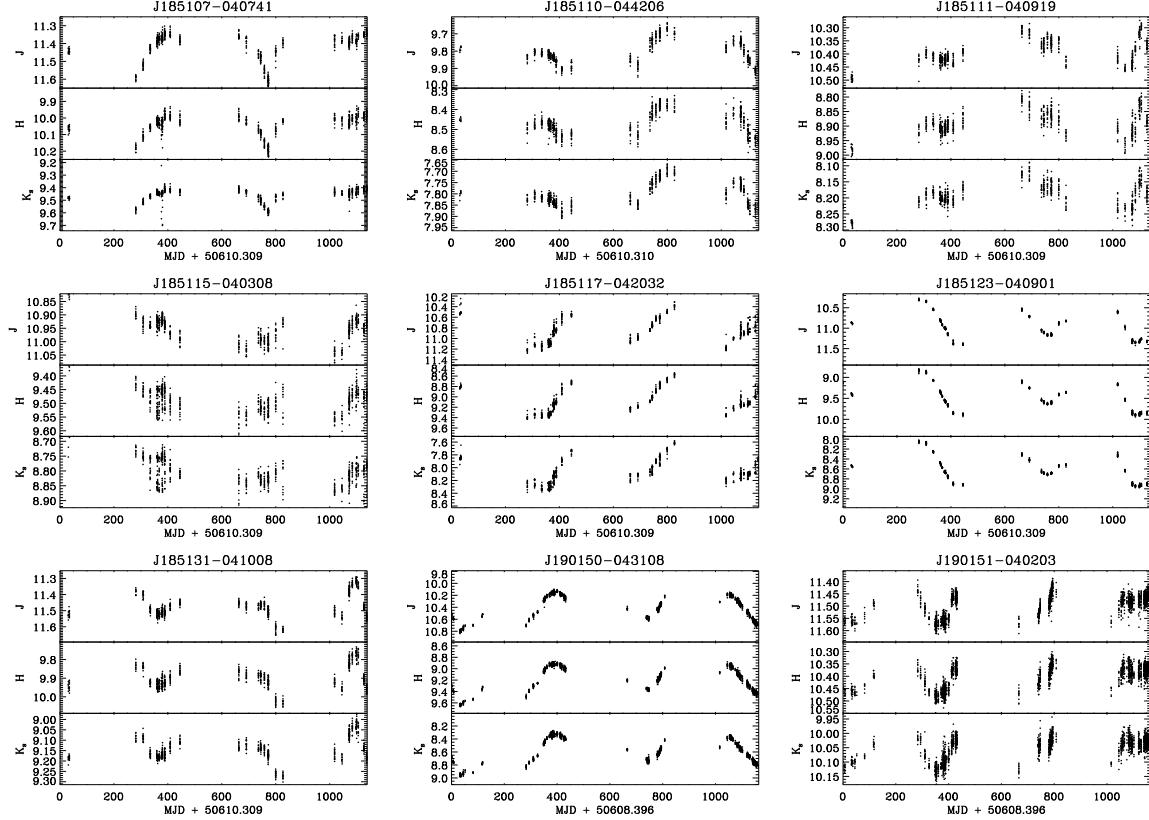


FIG. 16.— some long quasi-periodic variable objects. a couple look like DY Per's maybe? Nothing quite like RCB