# ASTR 511

## Galactic Astronomy

# Lecture 08

## Chemical Cartography

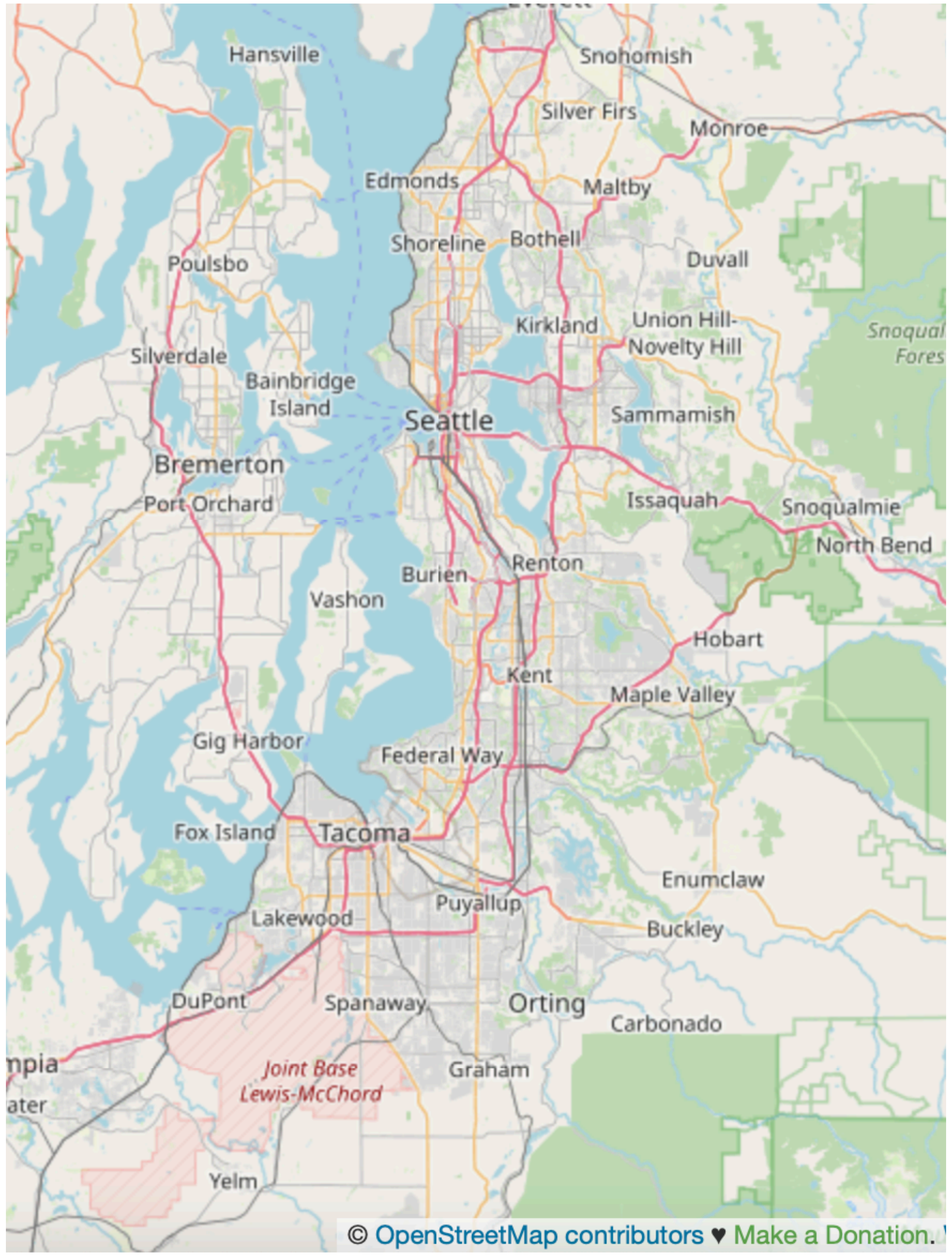
Prof. James Davenport (UW)

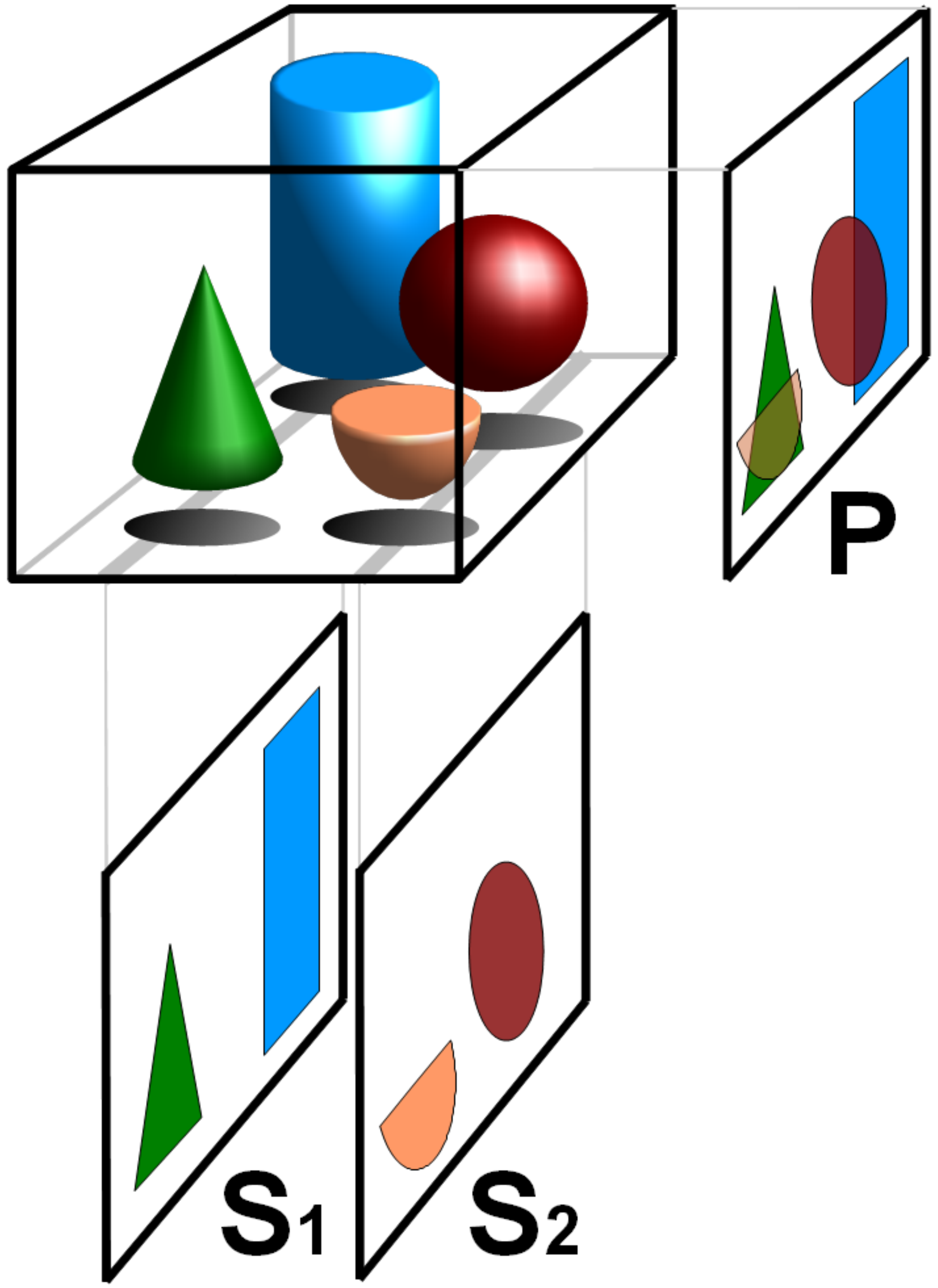
Winter 2023

# Cartography

*Tabula Rogeriana* (1154) Muhammad al-Idrisi

# Cartography versus Tomography




openstreetmap.org

https://commons.wikimedia.org/wiki/File:TomographyPrinciple_Illustration.png

# Both are tools in “Galactic Archeology”

# This is another shockingly deep field of study

i.e. we'd need *several* lectures to do it justice

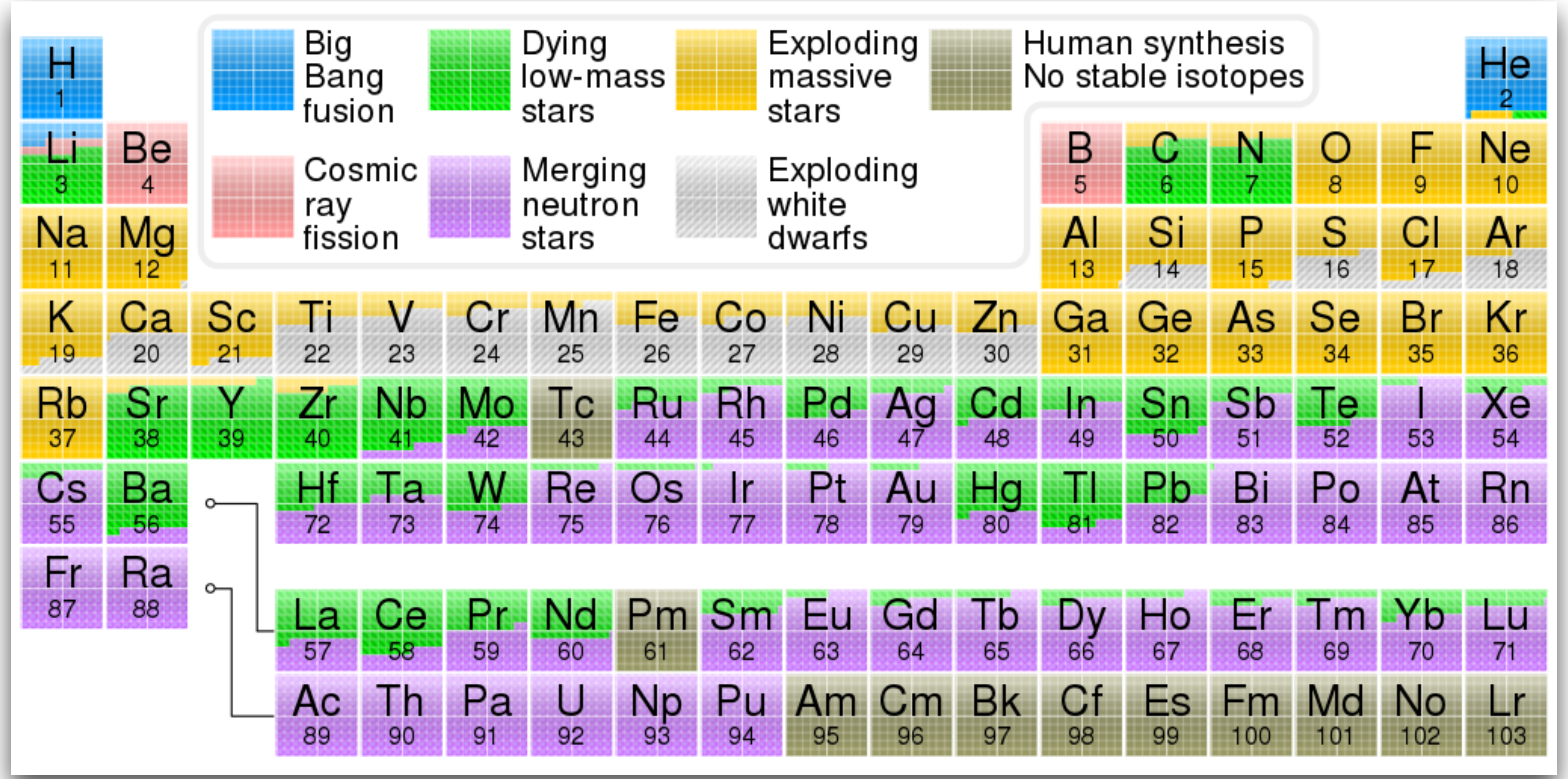
# Instead: we'll try to keep to 1 lecture & HW3

# Basic Story: Age -> Metallicity

- In the beginning there were no metals (Pop III stars)
- Then there were few metals (Pop II stars)
- Now there are increasingly more metals (Pop I stars)

- Stars burn H -> He, mess with other elements along the way (e.g. CNO)
- AGB stars generate s-process elements in shell layers of fusion
- SNe  quickly produce r-process elements

# Basic Story: Age -> Metallicity



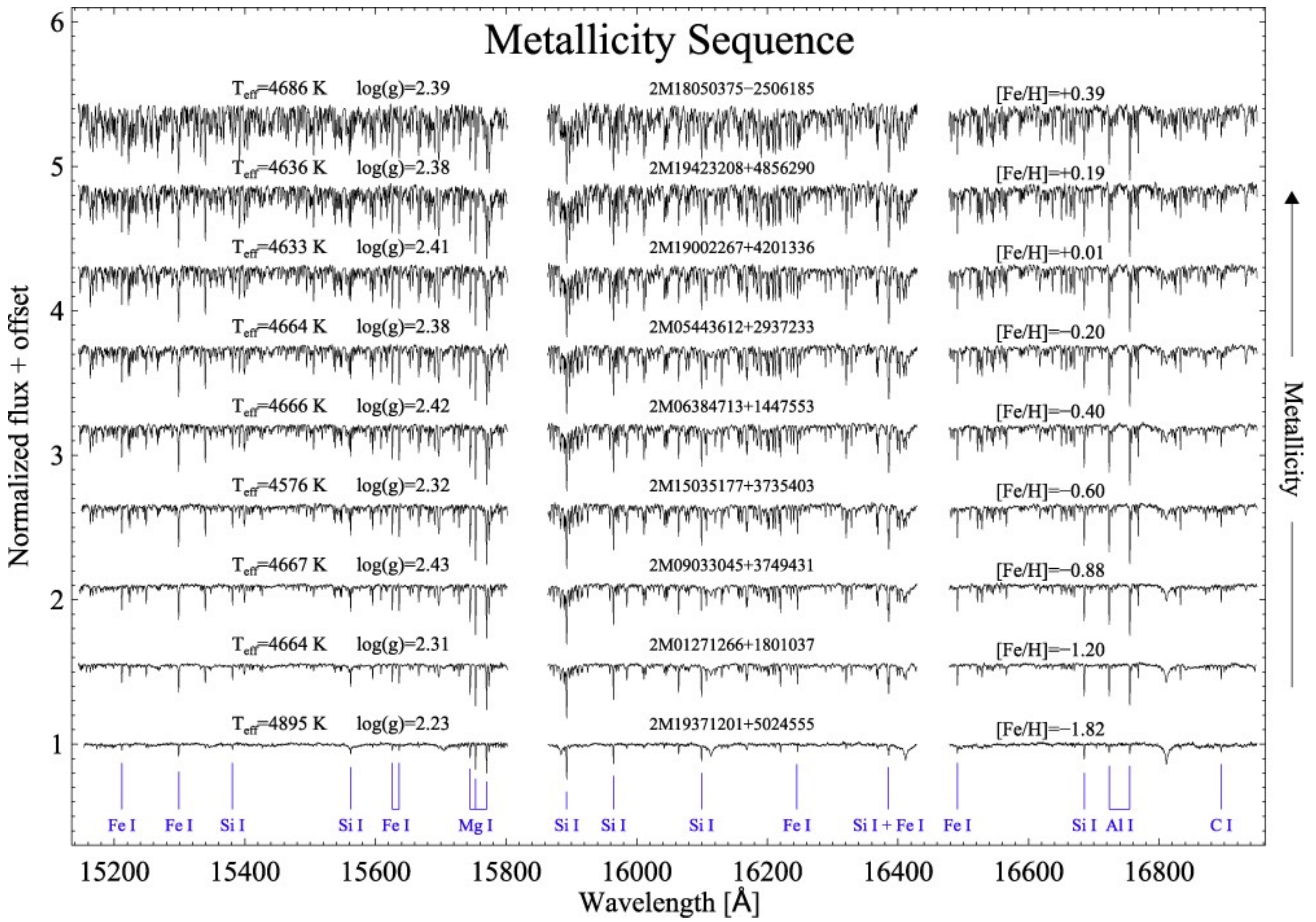

https://commons.wikimedia.org/wiki/File:Nucleosynthesis_periodic_table.svg

# Basic Story: Age -> Metallicity

- Metals, typically approximated as **[Fe/H]**, indicate age!
- As we've discussed, key to describing origin stories of every part of the Galaxy, young vs. old
- Affects stellar evolution (e.g. opacities, gas cooling, etc)
- Entire galaxies shown to have bulk metallicities, impact Star Formation, IMF

# Basic Story: Metallicity

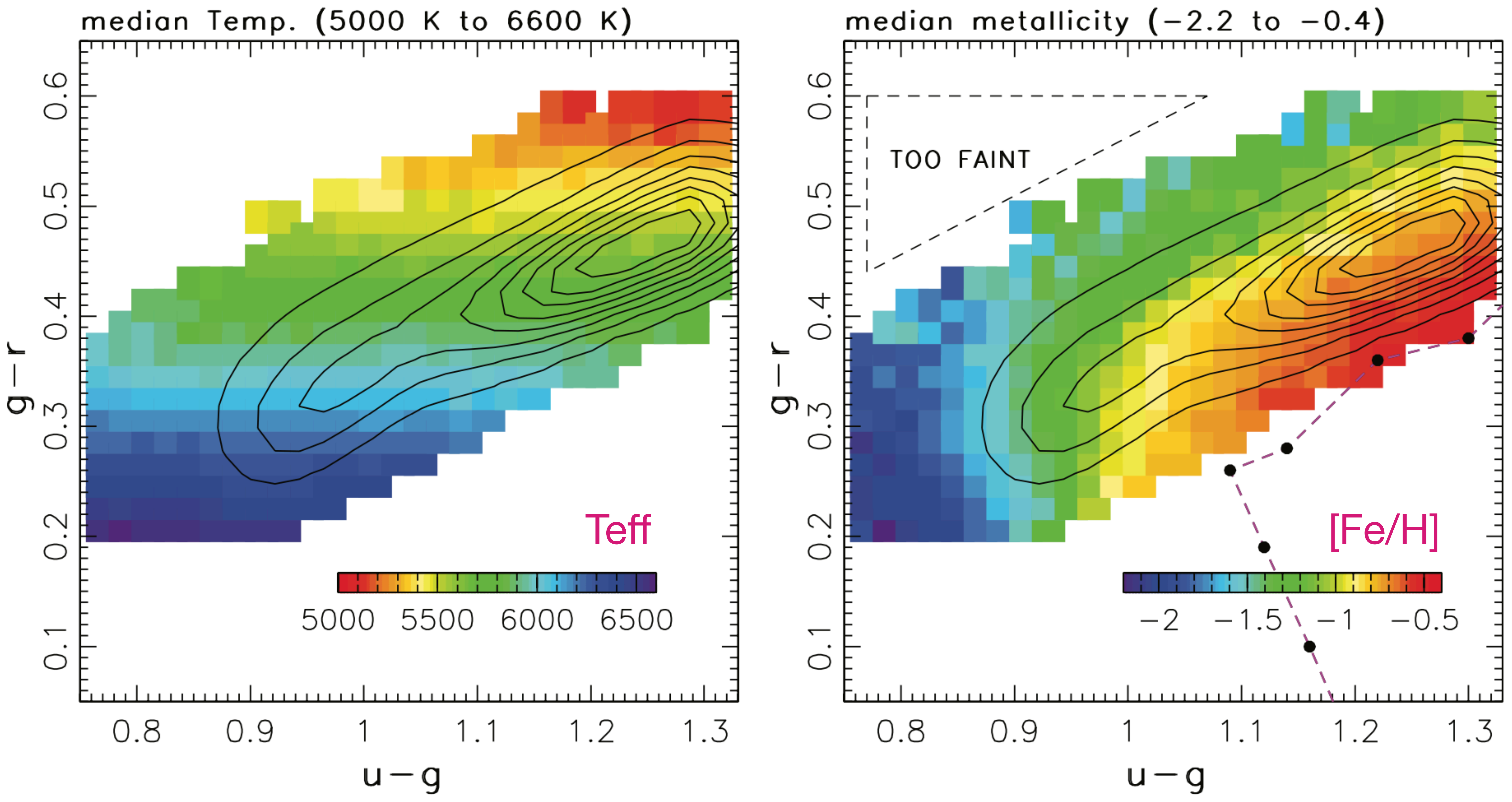
- Typically summed up as [Fe/H], i.e. the log ratio of Fe/H *relative* to the solar amount
    - Also abundances of individual elements are studied, as well as groups (e.g. [$\alpha$/Fe] )
- Primarily determined via spectroscopy, modeling atomic absorption lines
    - High resolution VERY helpful



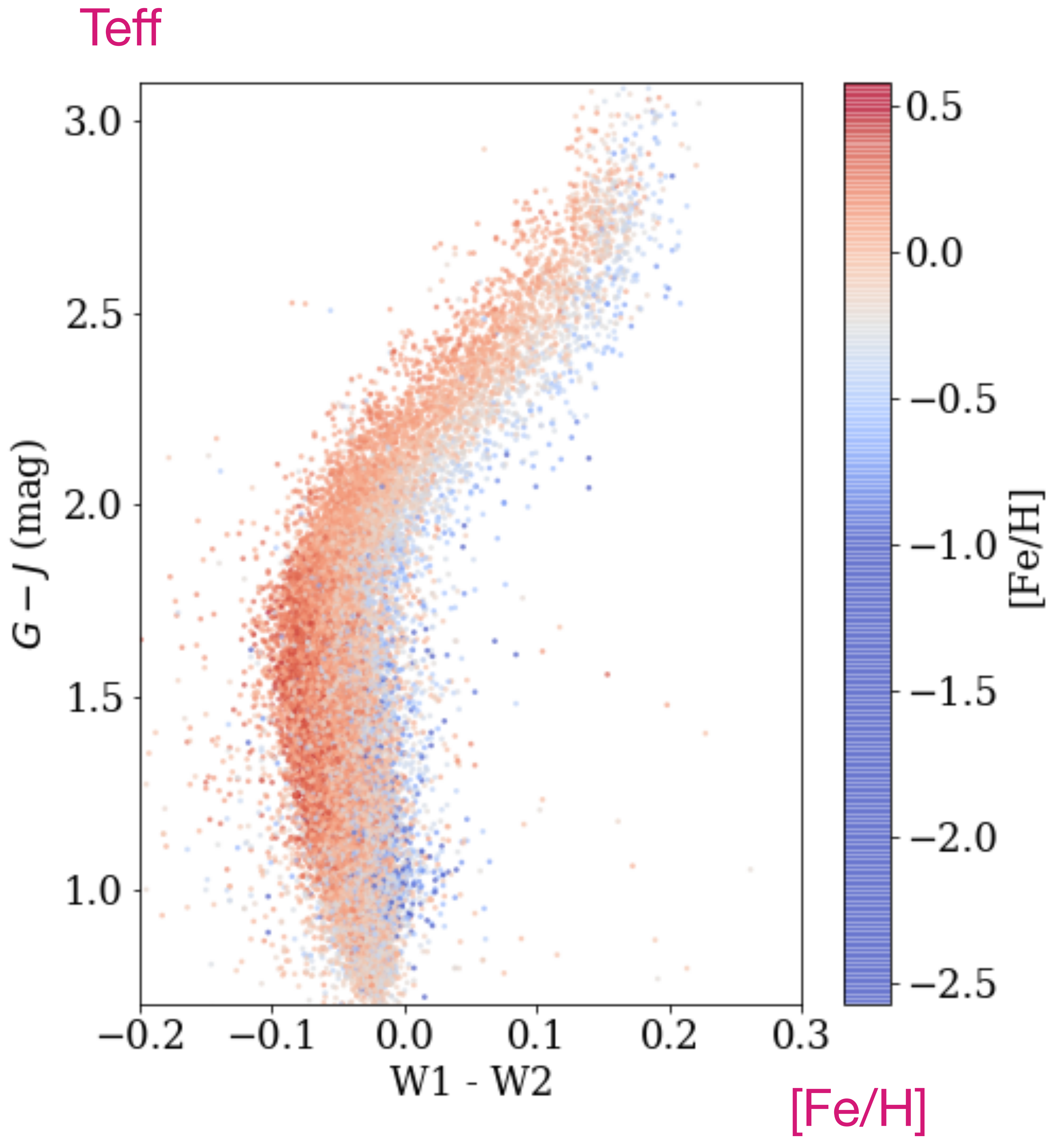

Majewski+2017 (APOGEE)

# Method 1. Photometric Metallicity

- Amazing for statistics with big surveys (hello: LSST!)




Ivezic+2008




Davenport & Dorn-Wallenstein (2019)

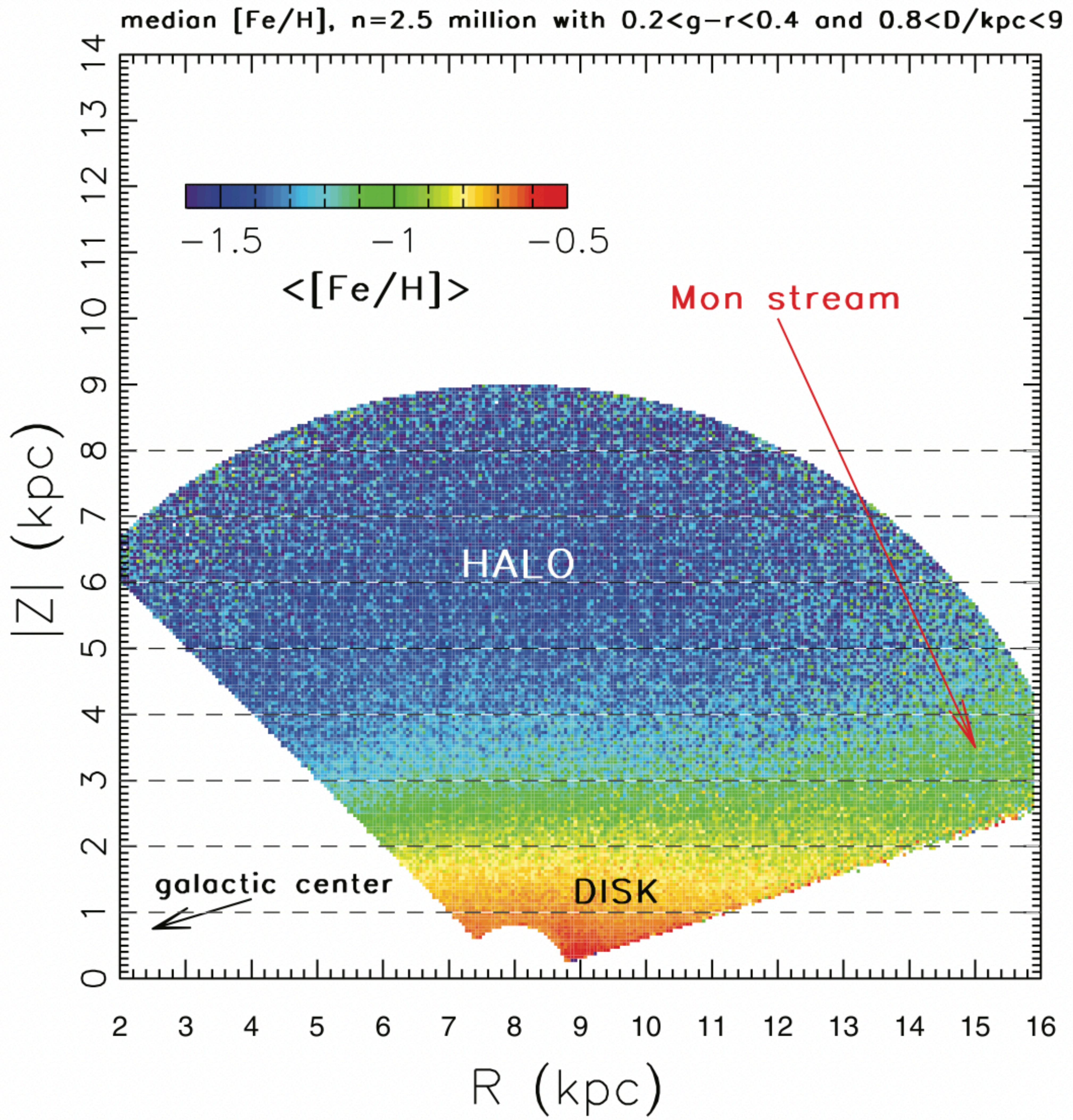
# Method 1. Photometric Metallicity

- Amazing for statistics with big surveys (hello: LSST!)

- Doing this for hundreds of thousands (or even millions) of stars enables new studies of the composition of our galaxy!

- Take slices: Galactic Tomography

- Wonderful new term: chemical cartography



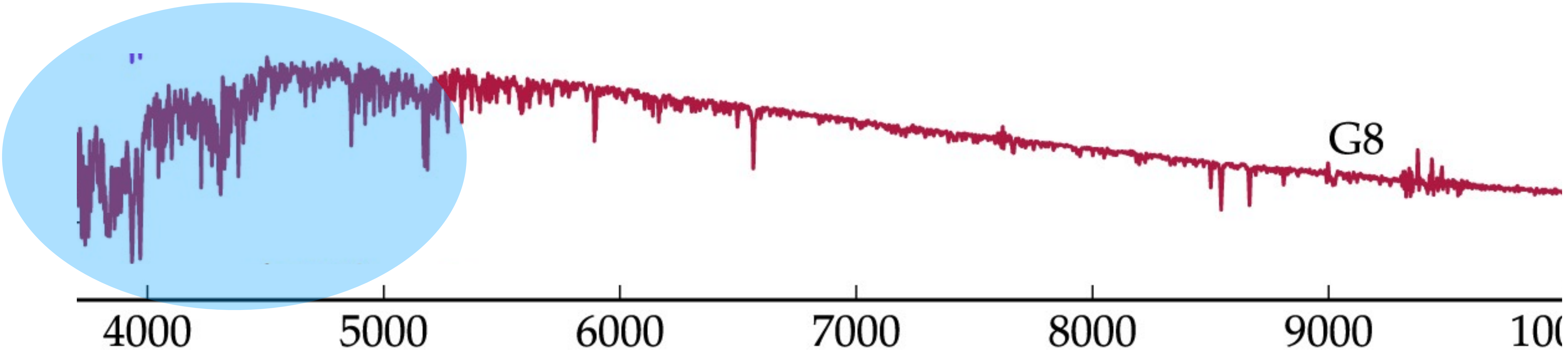

Ivezic+2008

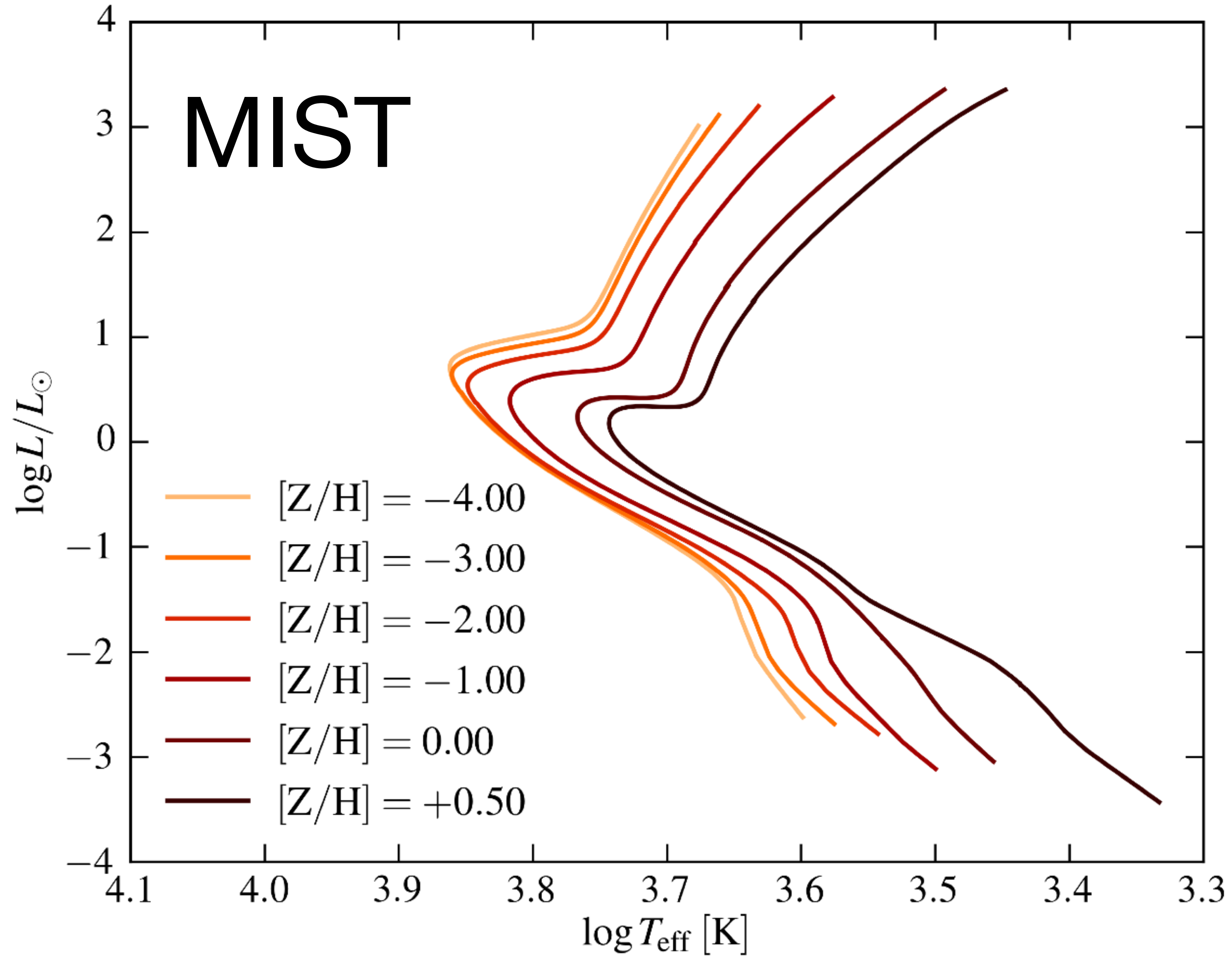
# Method 1. Photometric Metallicity

- Amazing for statistics with big surveys (hello: LSST!)

- In general, metal-poor: bluer (hotter $T_{eff}$) metal-rich, more lines, redder

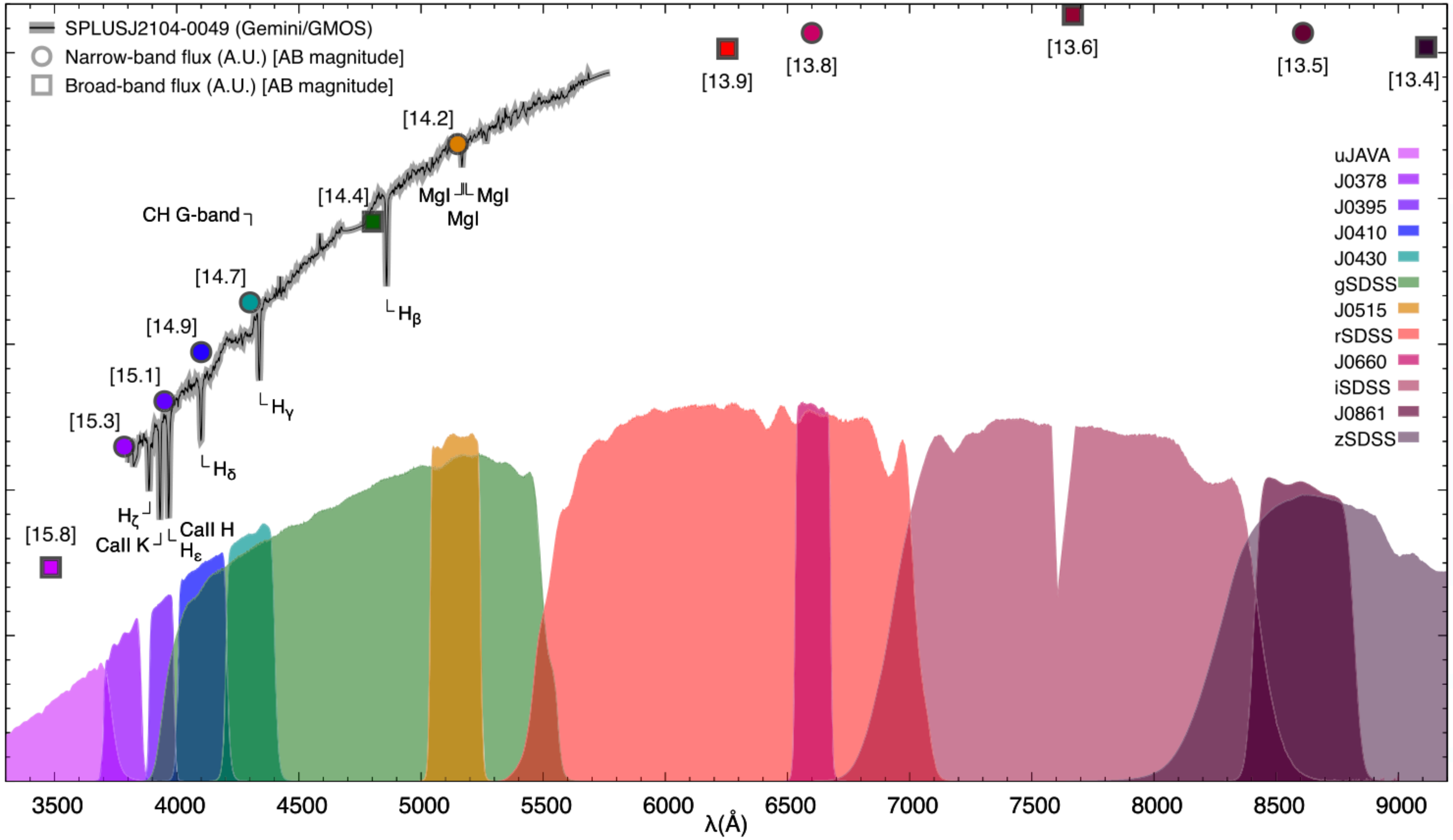
- Typically need blue (e.g. *u*-band) filters BUT, some sensitivity in the IR too



# Method 2. Narrow Band Filters

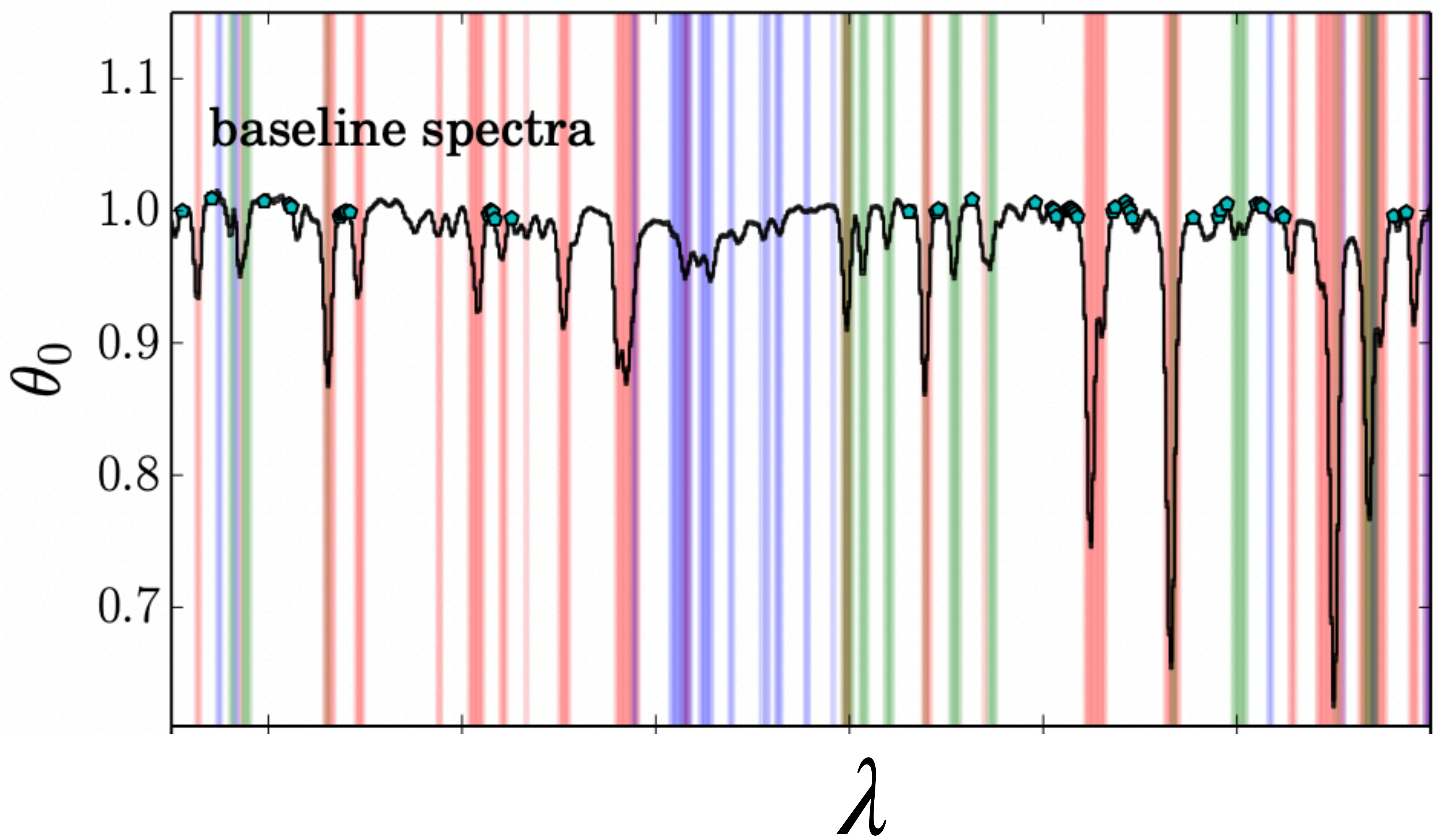
- Filters centered on spectral lines/features that are sensitive to specific elements
- Example: an "Ultra Metal-Poor Star" from S-PLUS
- **Not super popular (alas!) but long used/studied**
    - e.g. Strömgren filters!




Placco+2021

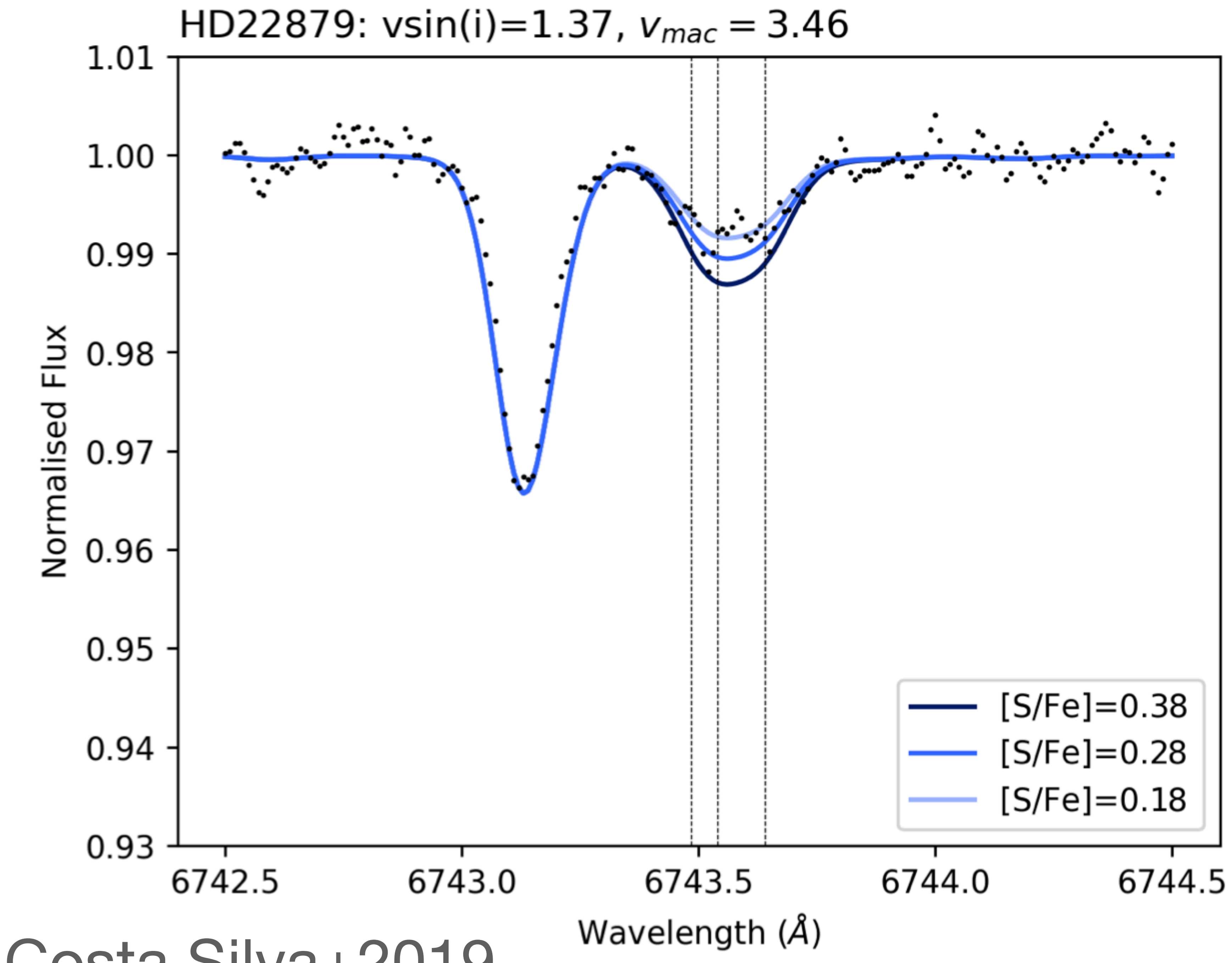
# Method 3: Spectroscopy

- Can be done with high or low resolution (higher-res better, obviously)
- Trace bulk metallicity [Fe/H] or [M/H], esp. with low-res fairly easily
- Track individual element species (beware: complex stellar atmosphere and "spectral synthesis" modeling challenges!)
- Some rad new data-driven tools (e.g. The Cannon)




[Fe/H]
log g
Teff

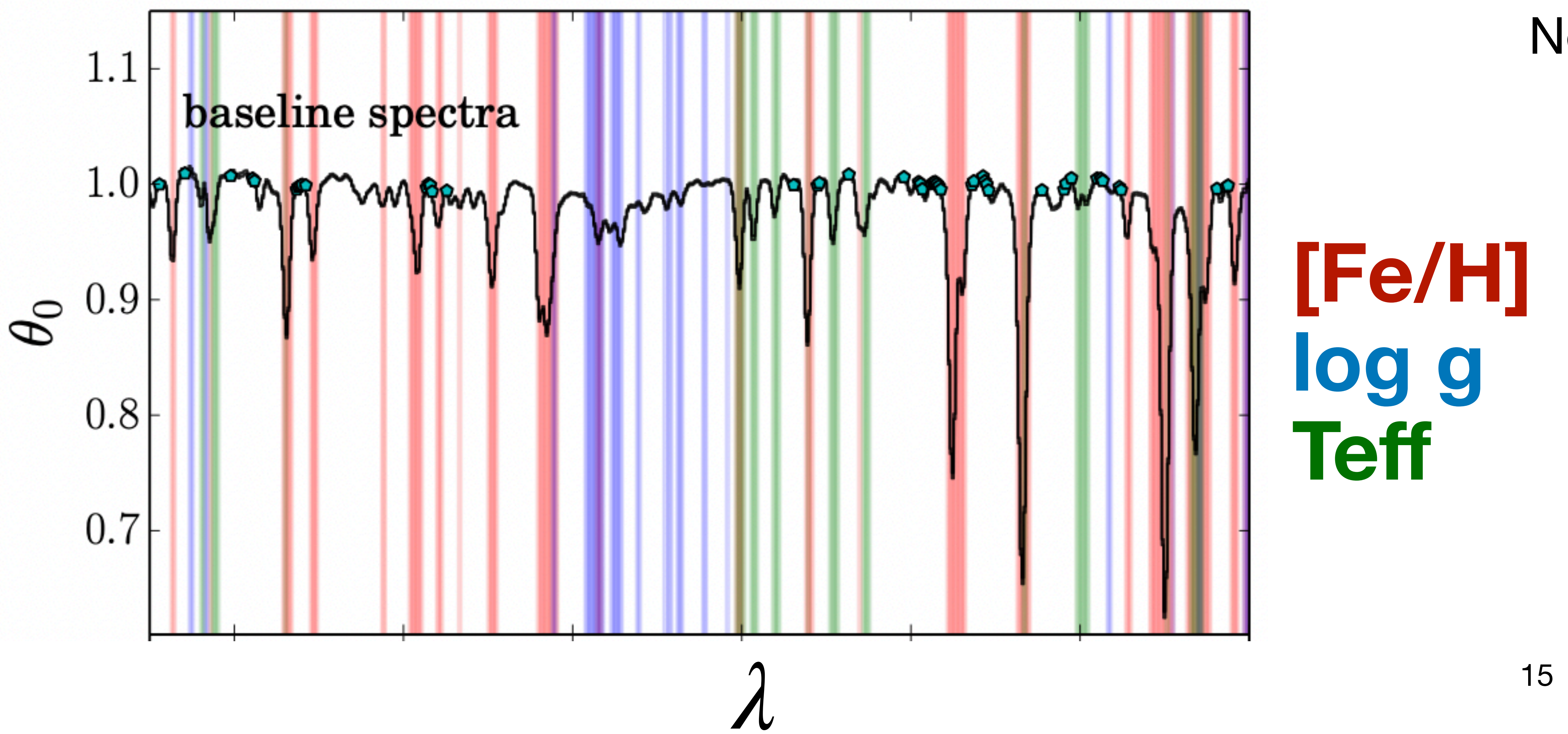
Ness+2015



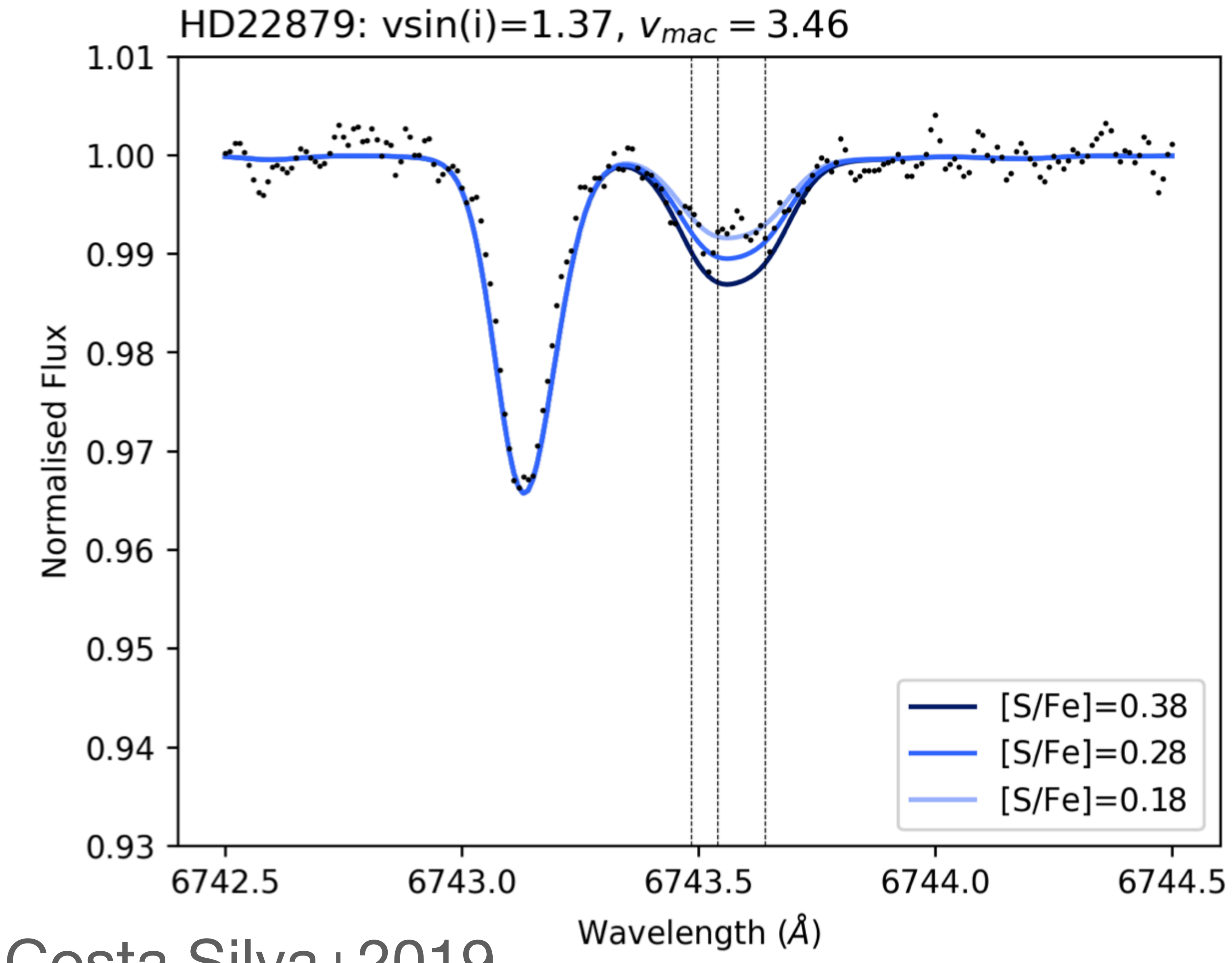

Costa Silva+2019

# Method 3: Spectroscopy

- We in the era of Spectroscopic Surveys!
  - SDSS (I-IV), APOGEE, GALAH, RAVE




Ness+2015




Costa Silva+2019

# Beyond [Fe/H] & Age

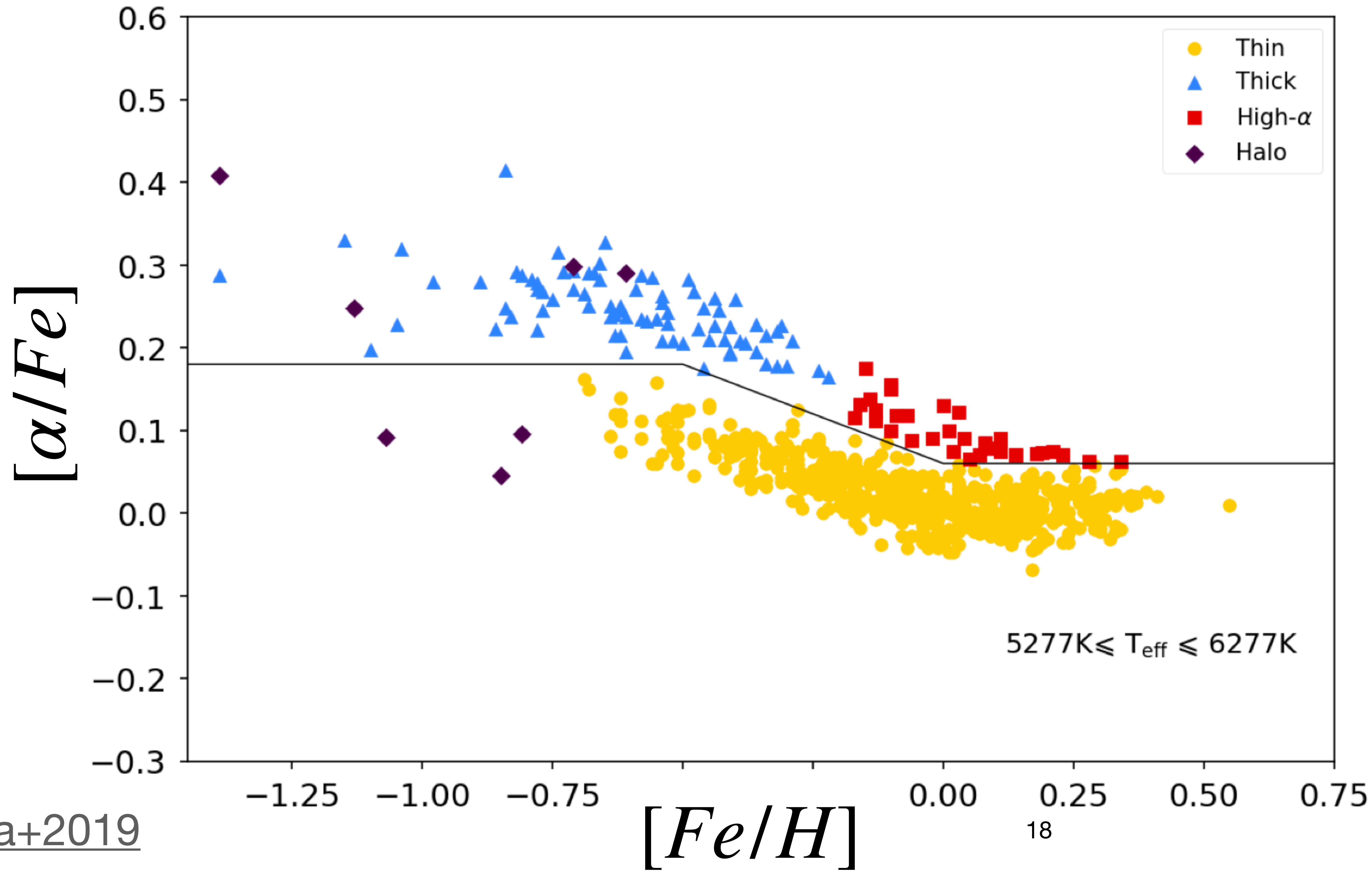
- Stars burn H -> He, mess with other elements along the way (e.g. CNO)
- AGB stars generate s-process elements in shell layers of fusion
- **SNe quickly produce r-process elements**
  - Not really the full story!

# Beyond [Fe/H] & Age

- Type II SNe (massive star) produce lots of alpha elements
  - e.g. Ne, Mg. Si, S, Ar, Ca, Ti, O
- Type Ia SNe (lower mass stars, WDs) produce “iron peak” elements (and also some alpha elements)
  - e.g. Mn, Fe, Co, Ni

# Alpha Elements

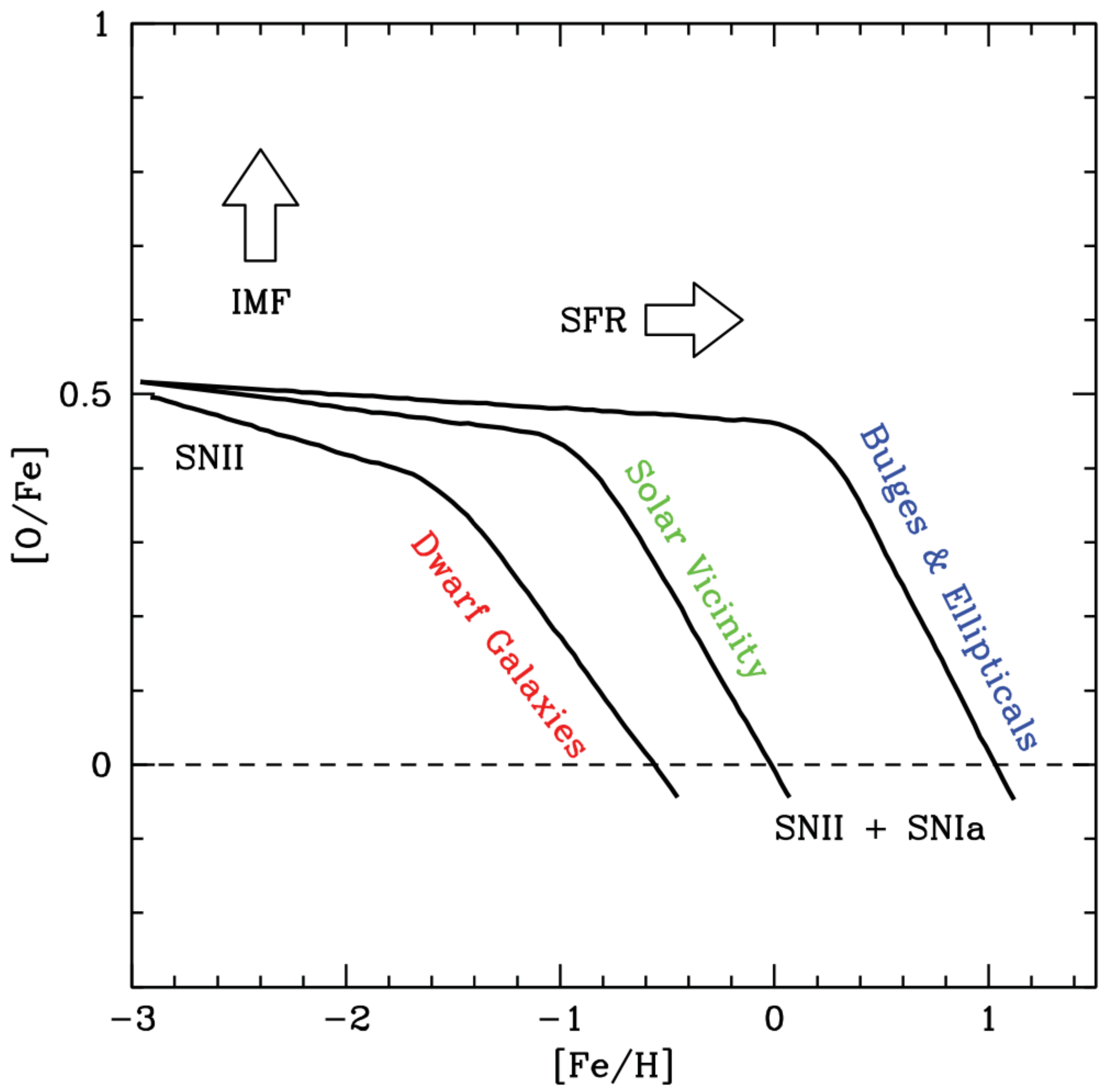
- Alpha elements / Iron Peak ≈ Massive SNe / Low-mass SNe
- Alpha elements critical for tracing evolution of galaxy!




Wallerstein–Tinsley Diagram

Costa Silva+2019

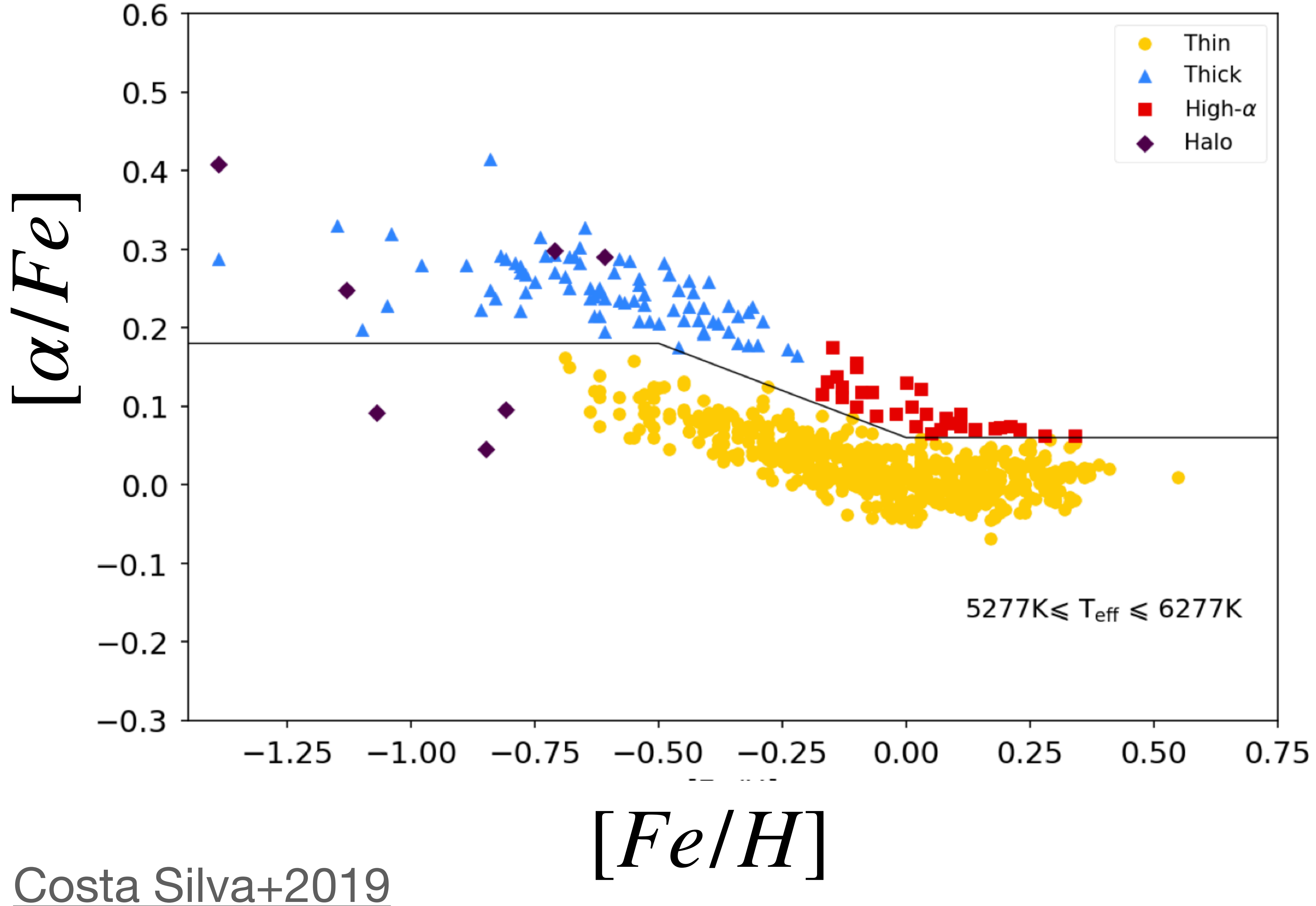
# Alpha Elements




McWilliam (2016), from
Matteucci & Brocato (1990)

# Alpha Elements

- **Thin & Thick disk have different star formation histories based on** $[\alpha/Fe]$




Costa Silva+2019

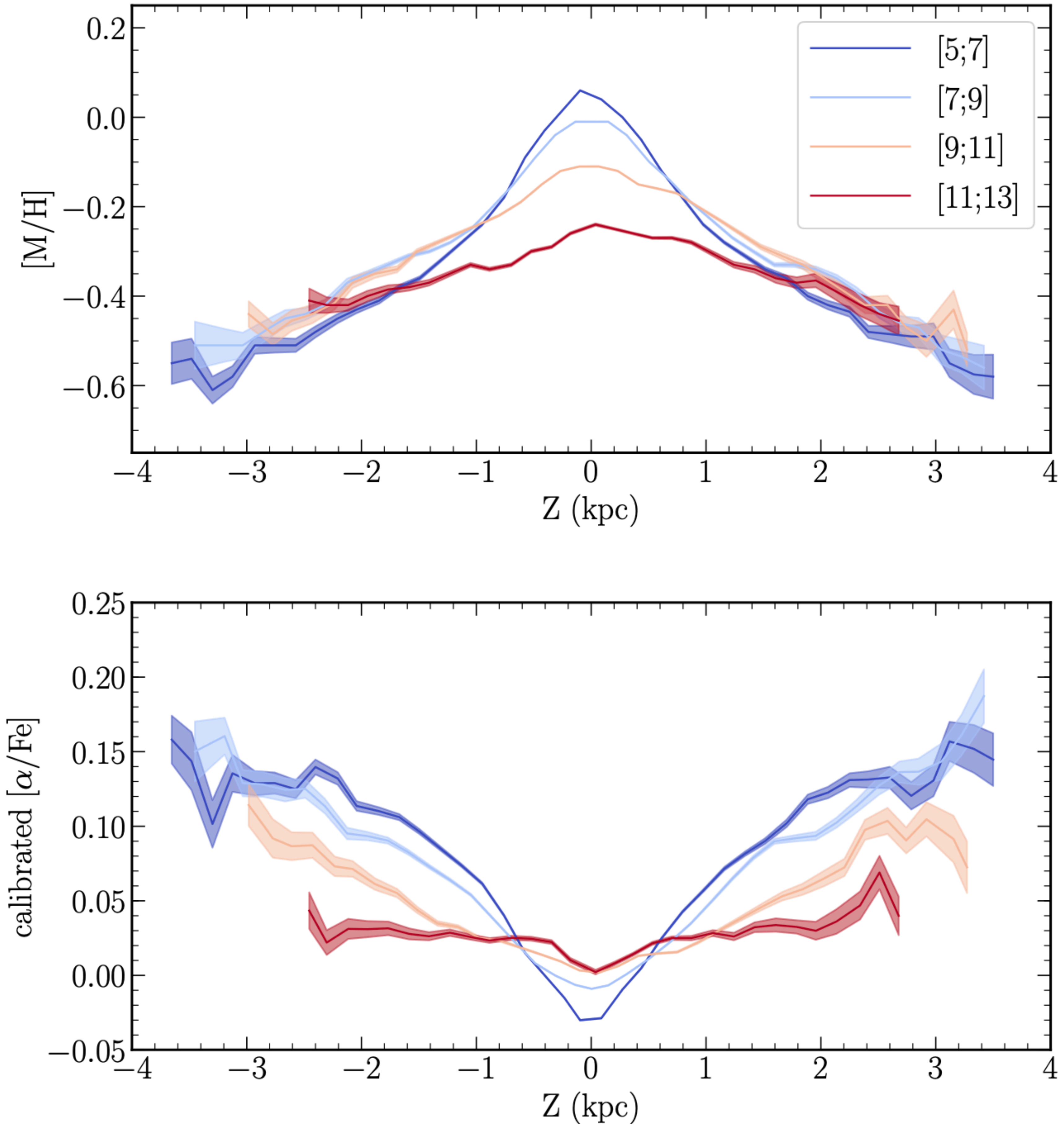
- Detailed Galactic Chemical Evolution models still tough
  - Review: Matteucci (2021) "…different chemical elements are produced on different timescales by stars of different masses."
- Lots of assumptions about enrichment timescales, "closed box" vs. in-fall, outflows, SFH…

# The "G Dwarf Problem"

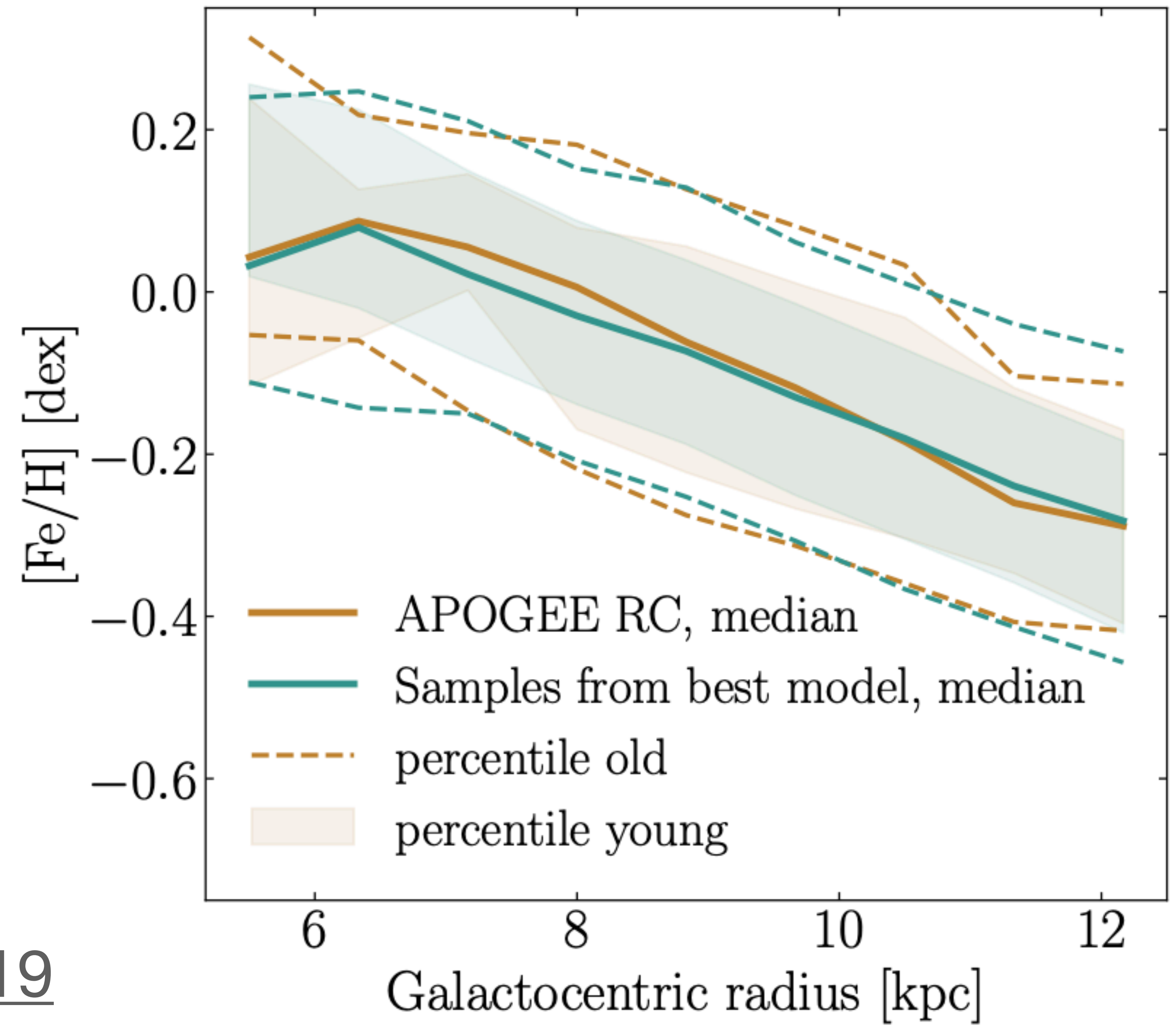
- A classical astronomy problem: If SFH is smooth/constant, and simple picture of age -> [Fe/H]… where are the low-metallicity G dwarfs?
    - Similar "problems" found for K dwarfs and M dwarfs
- **The simple model is broken** (for the MWY): we accrete gas, outflow of gas, IMF/SFR changes, varying mixing/recycling rates

# Inside-Out Growth of MWY Disk

- Metallicity lower in outer disk.
- Star formation more active in inner disk!

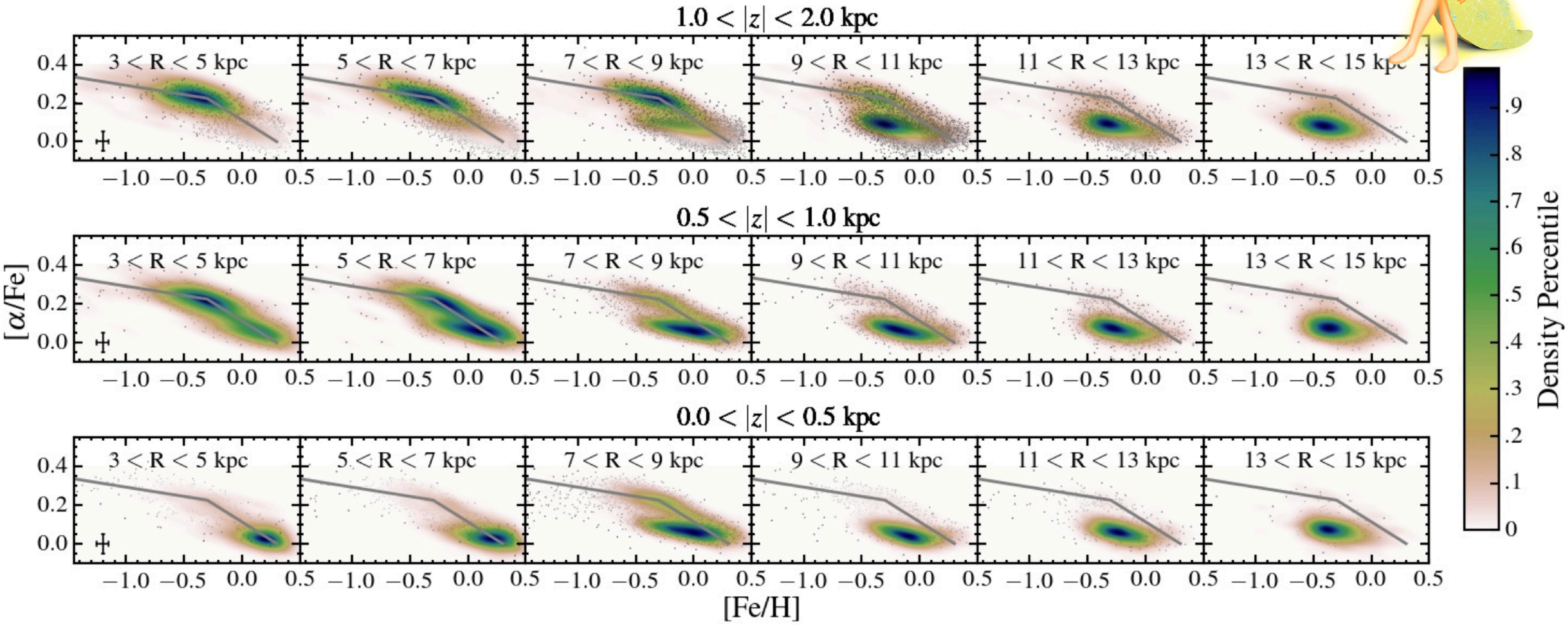



Gaia+2022




Frankel+2019

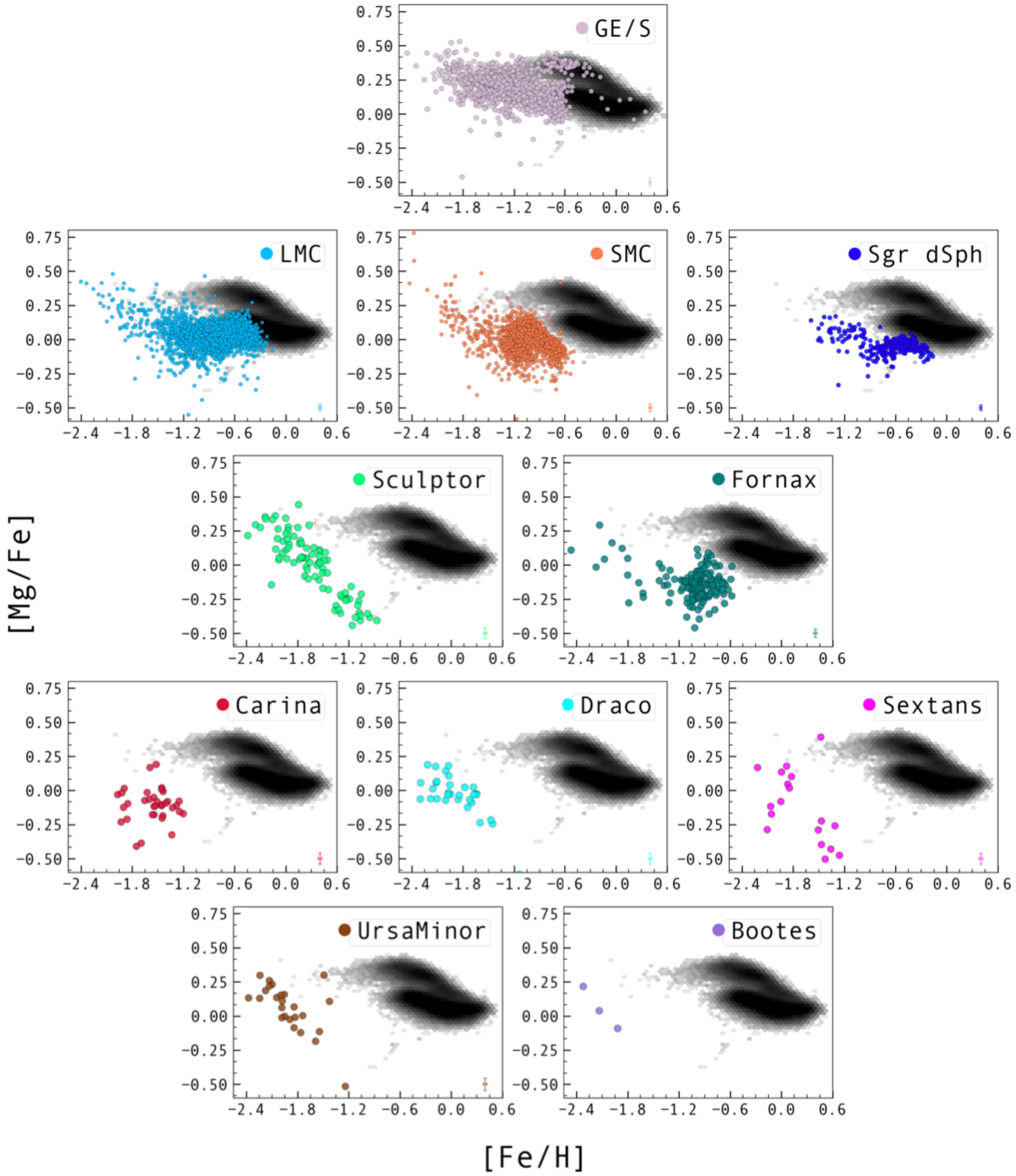
# Inside-Out Growth of MWY Disk




Hayden+2015

# Accretion History

- These accreted satellites (were: dwarf galaxies, now: tidal debris) have distinct chemistry from MWY
    - Also from SMC/LMC
-

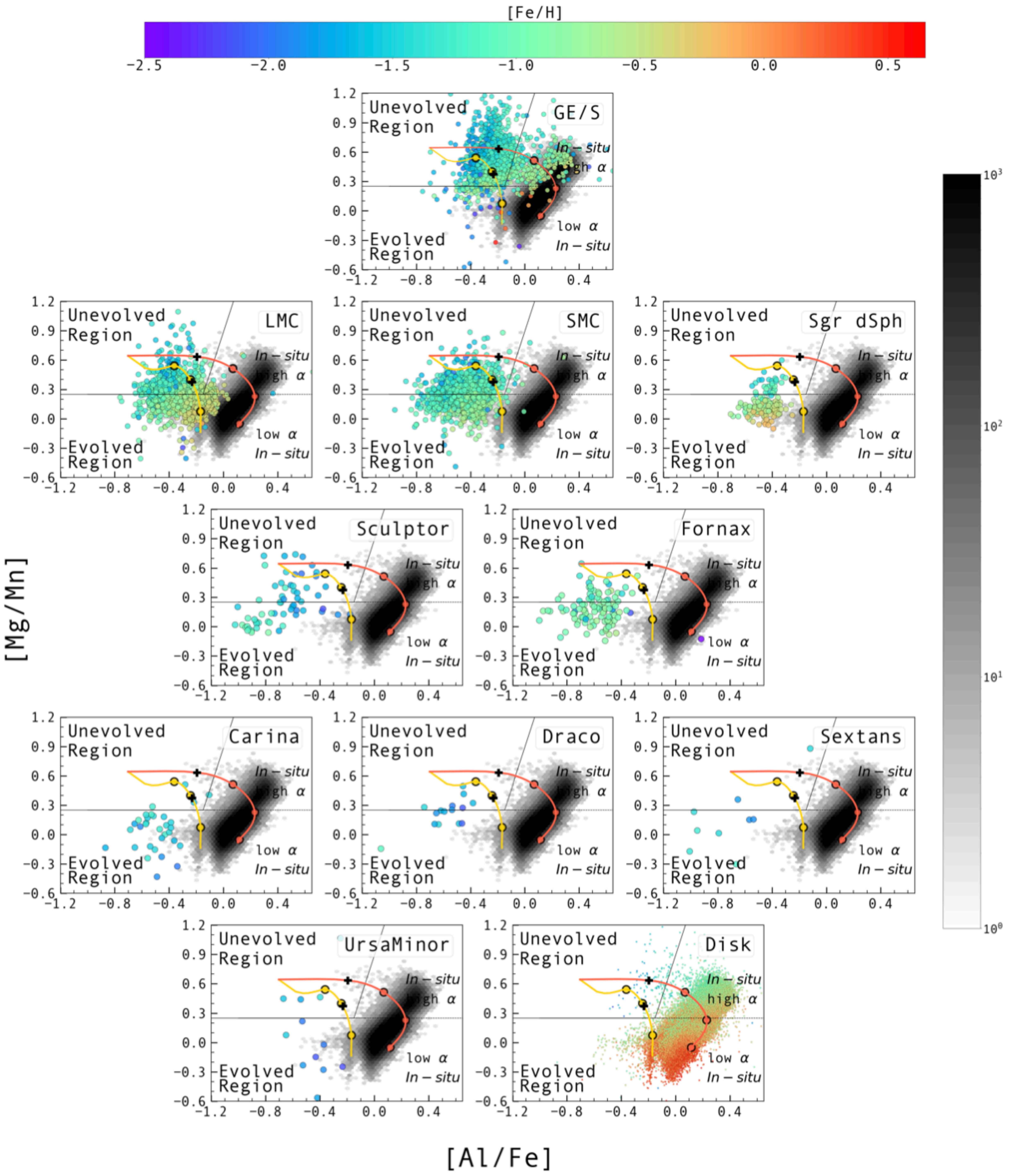



Fernandes+2023

# Accretion History

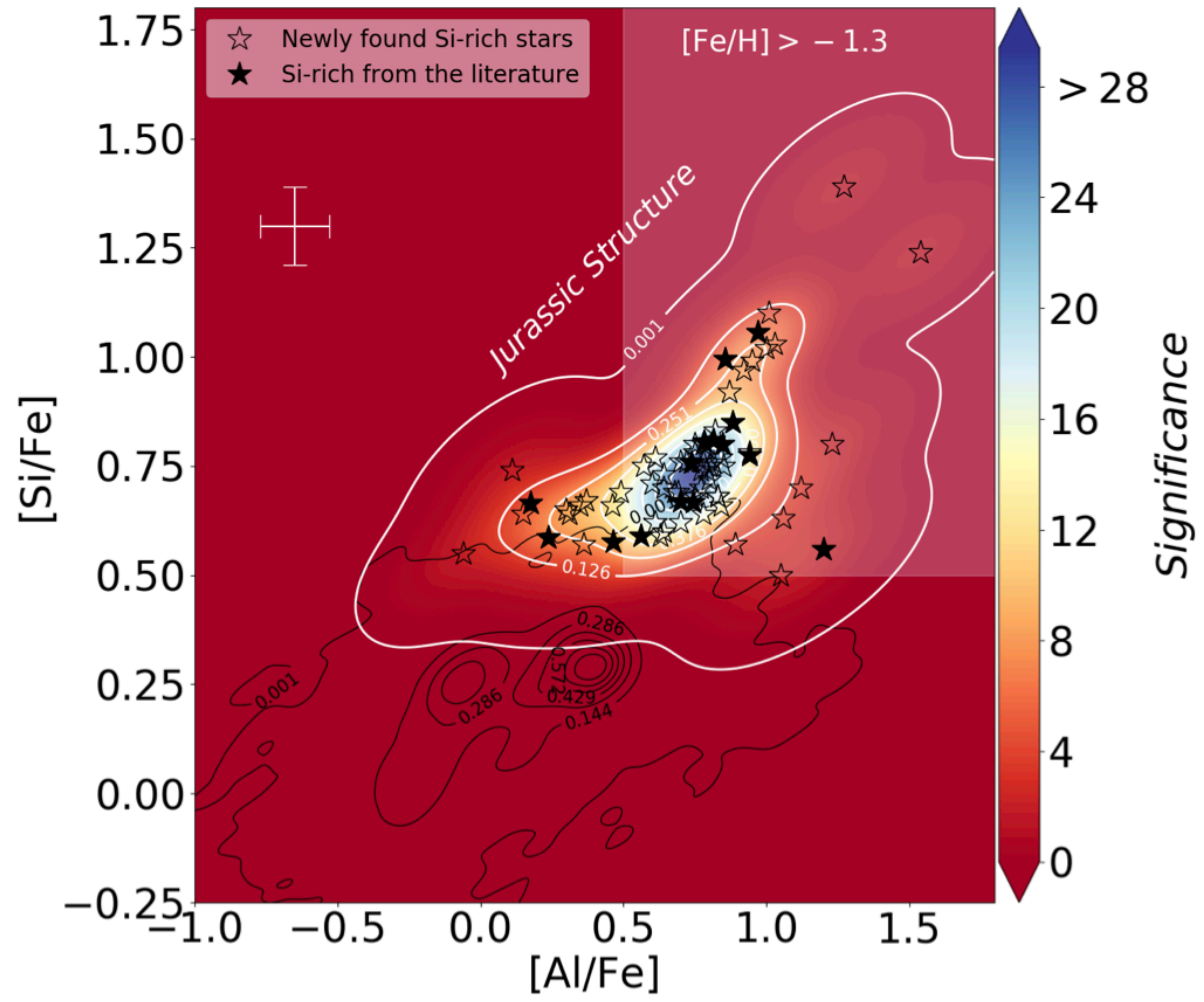
- Can be used to find NEW substructure (how much thick disk & halo come from distinct mergers?)
- Can help estimate mass of the merger
- Constraints on dSph formation, feedback, big simulations…

Fernandes+2023

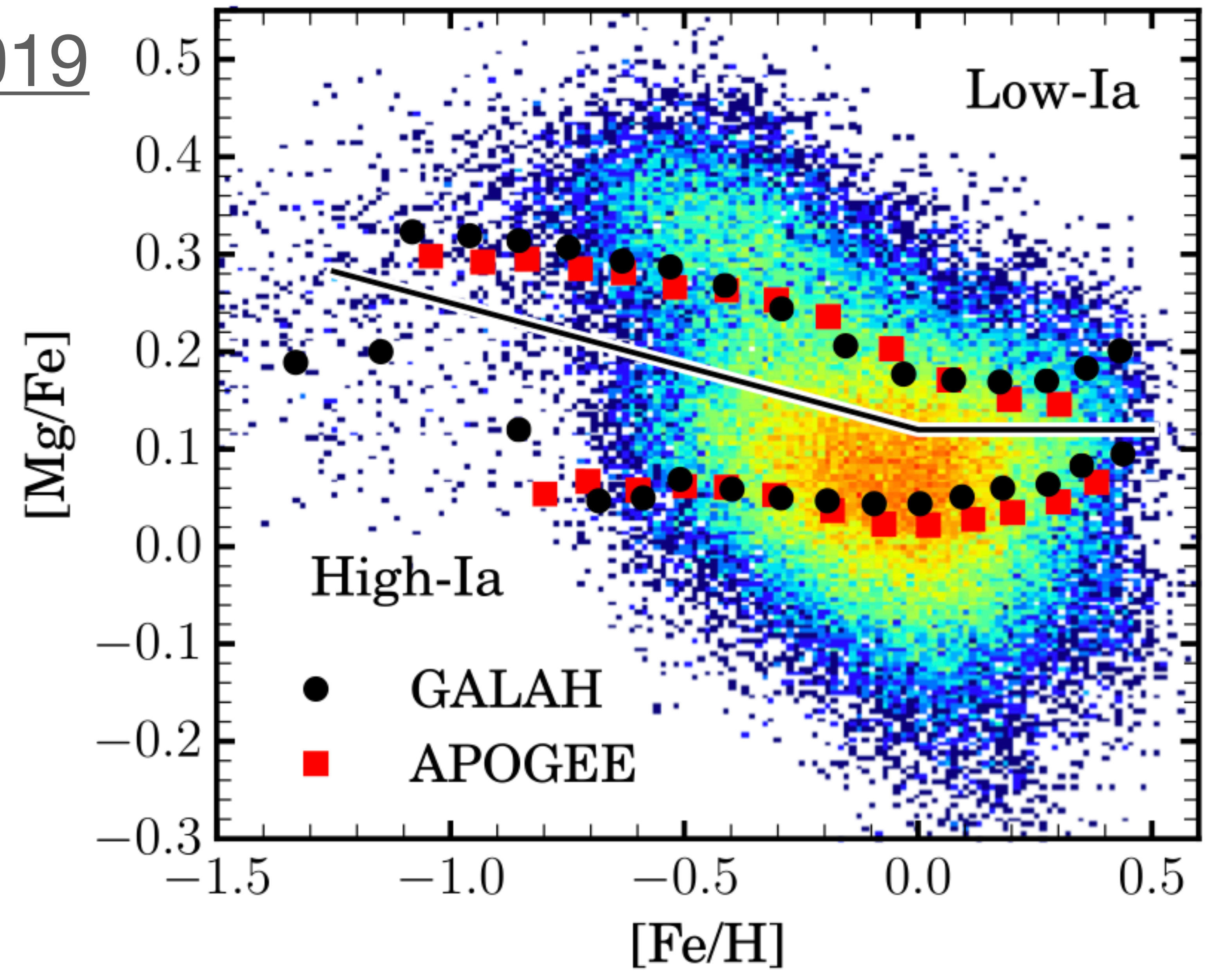


# Finding New Remnants/Structures

- GE/S is the big news, of course
- Many other distinct, smaller populations cropping up… with great names!
  - e.g. Sequoia, Jurassic



Griffith+2019

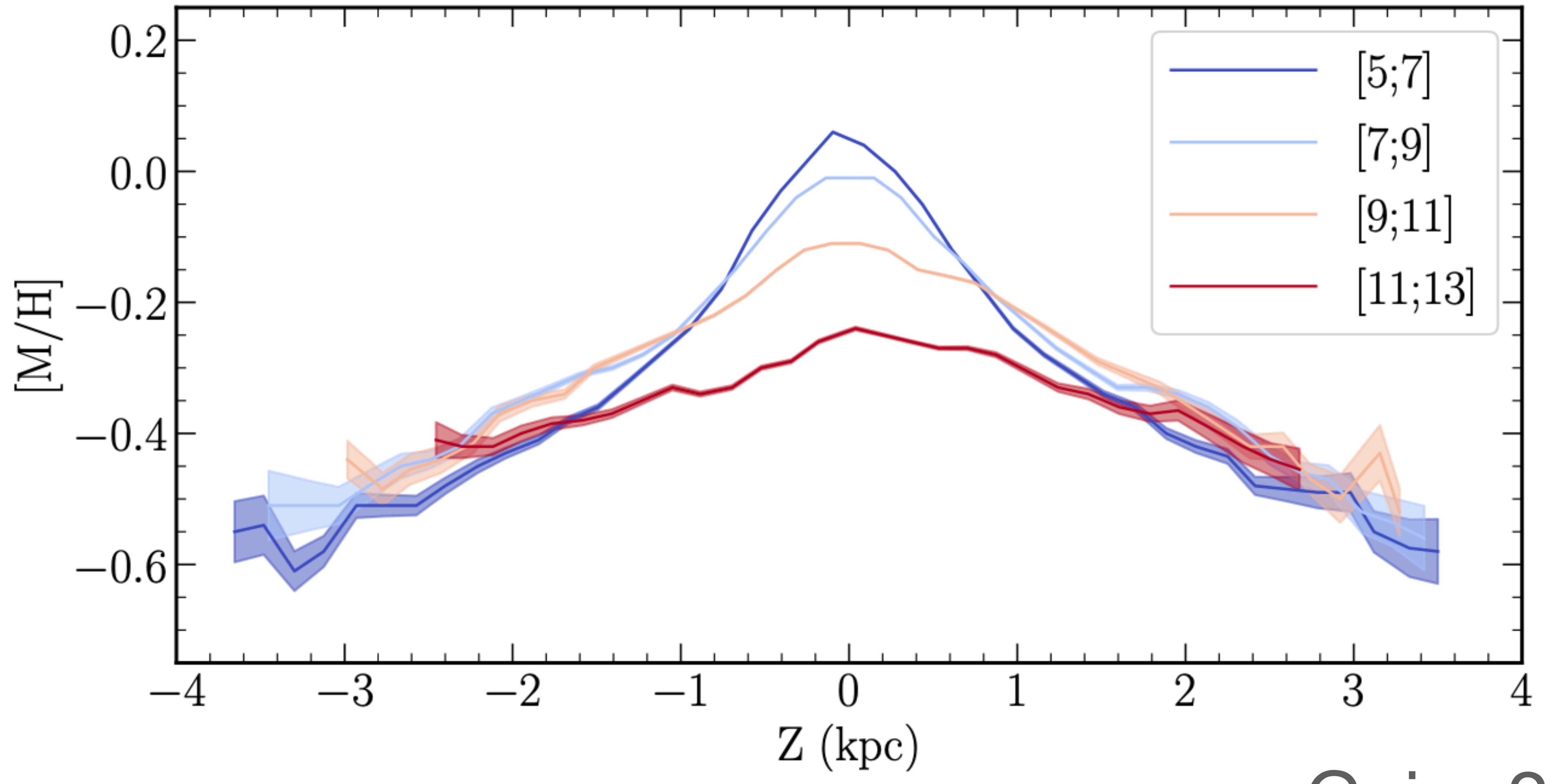


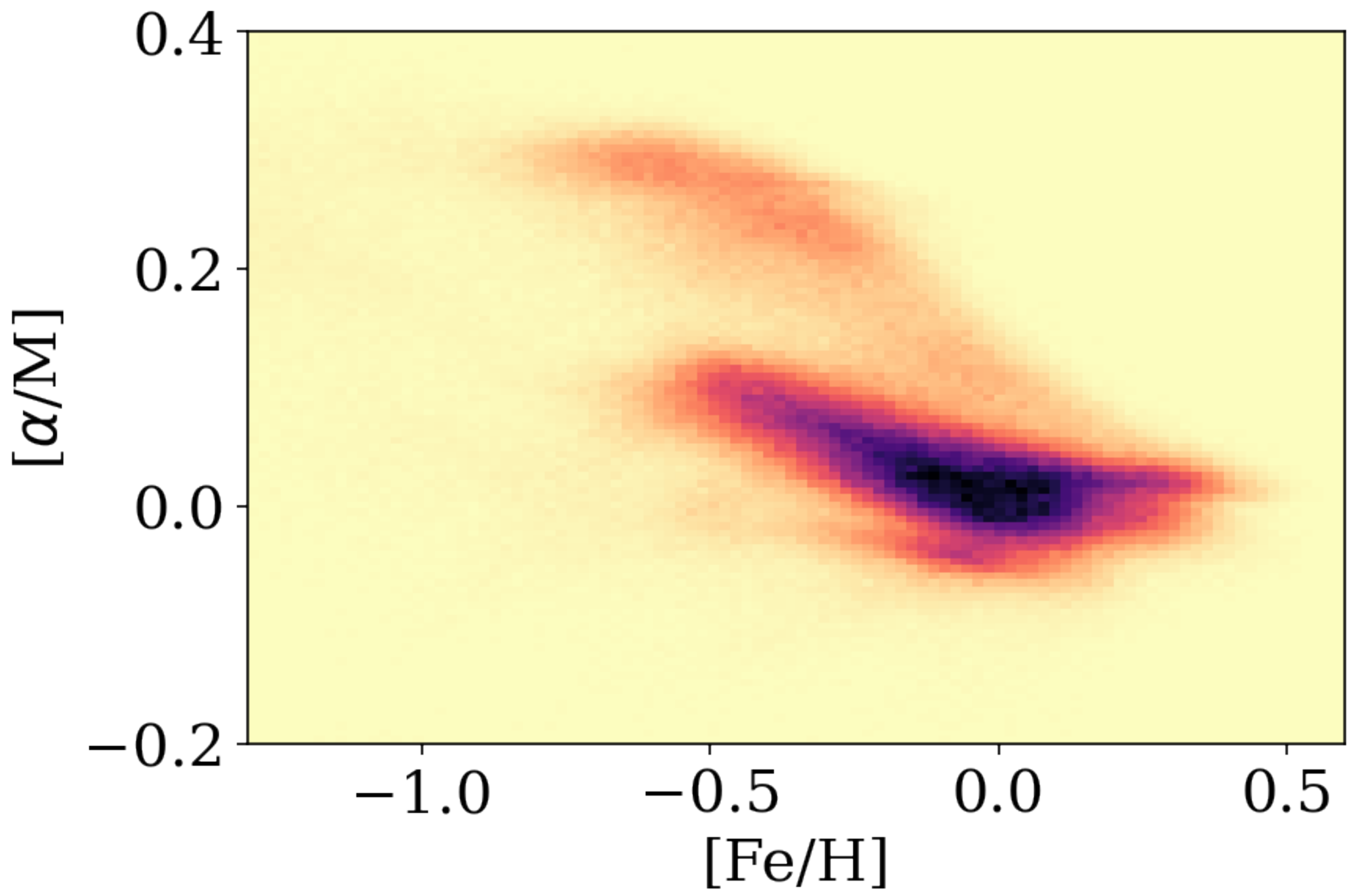

# *The Future*

- Finding more kinematic & chemical sub-structure will be productive for a while, more surveys contributing!
  - Probably some fun ML/algorithms stuff here
  - Starting to see time-domain (e.g. asteroseismology) play a part too
- Improved galactic chemical evolution models,
  - Model star formation within disk(s)/halo/bulge + *many* specific mergers
  - Detailing specific element ratios beyond simply $Fe$, $\alpha$, etc
  - Improved enrichment and mixing timescales (globular cluster mult. pops?)

# Homework 3

- **Look at the Wallerstein – Tinsley diagram from APOGEE**
- Thin/thick disk are easily visible
  - A *bunch* of sub-populations too!
- Disk structure!




Gaia+2022

# Next Time

- Star Formation History
  - How do we measure stellar ages??