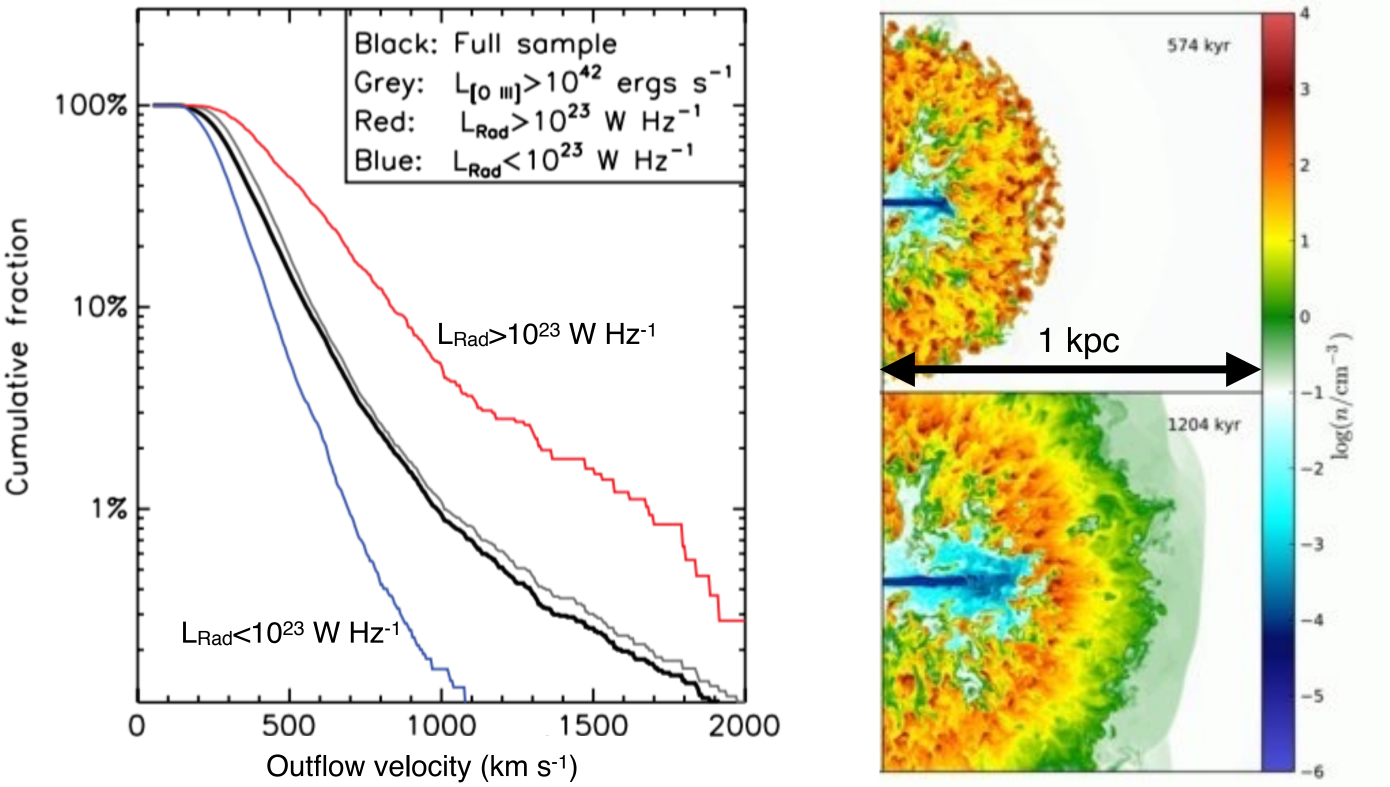

Fig. 1 *Left:* The fraction of emission line-selected AGNs in the SDSS at z<0.4 with (unresolved) outflow velocities greater than the value shown on the x-axis (from Mullaney et al. 2013). Outflow velocities are measured from the profile of the [O III]5007 emission line. The sample has been split according to radio luminosity and [O III] luminosity (see labels). A much higher fraction of powerful radio AGNs display evidence of fast outflows compared to low radio luminosity AGNs. This indicates that radio luminous AGNs are more capable of driving outflows, yet most AGN IFU surveys have focussed on radio-weak AGNs. *Right:* Simulated interaction between an AGN-launched radio jet and the interstellar material, 500kyr and 1.2 Myr after jet onset (from Wagner et al. 2011). These hydrodynamical simulations demonstrate the capability of radio jets in disrupting dense gas within the host galaxy. Our MUSE observations will map-out the kinematics of the outflows in radio AGN to determine whether they are, indeed, jet driven.

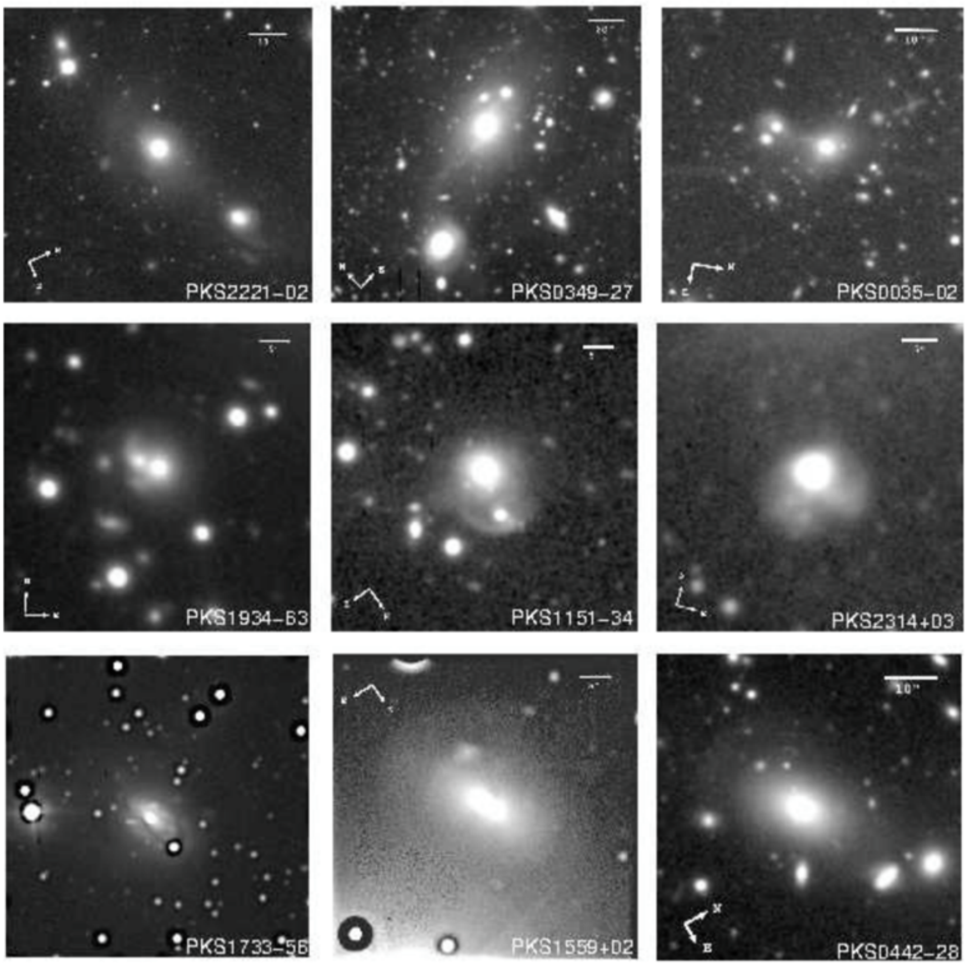

Fig. 2 Examples of deep optical r-band images of 2 Jy radio galaxies from Ramos Almeida et al. (2011). The angular scale is shown by the white line in the upper right of each panel which is 5” long. While all are hosted in early-type galaxies, they display significant diversity in their detailed structures. While PKS2221-02, PKS0349-27, PKS0035-02, PKS1934-63 and PKS1151-34 are pre-coalescence systems that show evidence for tidal interactions with neighbouring galaxies, PKS2314-03, PKS1733-56, PKS1559+02 and PKS0442-28 are post-coalescence systems that show evidence for high surface brightness tidal features (PKS2314+03, PKS1733-56) or more subtle shell structures (PKS1559+02, PKS0442-28). These complex structures indicate that radio galaxies have a far more violent recent merger history than their early-type morphologies initially suggest. However, it is not clear what types of interactions are most likely to trigger a radio AGN.