The background of the slide is a stylized, artistic representation of gravitational waves. It features a series of concentric, swirling lines in shades of blue and black, creating a sense of depth and motion. In the center, two yellowish-green spheres are depicted, representing the merging of two black holes. The overall effect is a dynamic and scientific illustration of the phenomenon being discussed.

Coherent and Coincident Analyses of LIGO–Virgo Data from the Third Observing Run

Candidate: Giada Caneva Santoro
Supervisor: Dr. Francesco Pannarale
Sapienza University of Rome

Gravity: from apples to ripples

GENERATION OF SPACETIME METRIC WAVES

Newton's law of
universal gravitation

$$F = G \frac{m_1 \times m_2}{d^2}$$

Einstein field equations
(EFE)

$$G_{\mu\nu} = \frac{8\pi G_N}{c^4} T_{\mu\nu}$$

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Weak field regime EFE
(outside source)

$$\square \bar{h}_{\mu\nu} = 0$$

**WAVE EQUATION WITH
PROPAGATION VELOCITY c**

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By imposing gauge conditions its 10 components reduce to two degrees of freedom, thus for a plane wave along the z axis:

$$h_{\mu\nu}^{TT} = \begin{bmatrix} 0 & 0 & 0 & 0 \\ 0 & h_+ & h_\times & 0 \\ 0 & h_\times & -h_+ & 0 \\ 0 & 0 & 0 & 0 \end{bmatrix} \cos \omega(t - z/c)$$

h_+ and h_\times are the two GW polarisations perpendicular to the direction of propagation

Gravitational Wave Detection

GRAVITATIONAL PERTURBATION ON MATTER

A GW generates periodic distortions in flat spacetime, which can be described in terms of the Riemann tensor:

$$R^\mu_{\nu\rho\sigma} = \frac{1}{2}(\partial_\rho\partial_\nu h^\mu_{\nu\rho} + \partial_\sigma\partial^\mu h^{\nu\rho} - \partial_\rho\partial^\mu h^{\nu\sigma} - \partial_\sigma\partial_\nu h^\mu_\rho)$$

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- Embodies the tidal force field and describes the relative acceleration between two particles in free fall.

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- Measures the curvature of spacetime.
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How does the curvature of spacetime affect the motion of matter?

Geodesic deviation: relative motion of nearby particles in terms of a tidal force determined by the Riemann tensor.

$$\frac{D^2\zeta^\mu}{D\tau^2} = -R^\mu_{\nu\rho\sigma}\zeta^\rho\frac{dx^\nu}{d\tau}\frac{dx^\sigma}{d\tau}$$

Gravitational Wave Detection

GRAVITATIONAL PERTURBATION ON MATTER

Geodesic deviation for two test masses in a near local-Lorentz frame, moving non-relativistically:

$$\ddot{\zeta}^j = -R_{0j0}^i \zeta^j$$

$$R_{0j0}^i = R_{i0j0} = -\frac{1}{2}\ddot{h}_{ij}^{TT}$$

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GW travelling in the z-direction towards a ring of test masses:

- The plus polarisation is

$$h_{ab}^{TT} = \begin{bmatrix} 1 & 0 \\ 0 & -1 \end{bmatrix} h_+ \cos \omega t$$

- The displacement between geodesic is

$$\zeta_a(t) = (x_0 + \delta x(t), y_0 + \delta y(t))$$

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$$\delta \ddot{x} = -\frac{h_+}{2} (x_0 + \delta x) \omega^2 \sin \omega t$$

$$\delta \ddot{y} = \frac{h_+}{2} (y_0 + \delta y) \omega^2 \sin \omega t$$

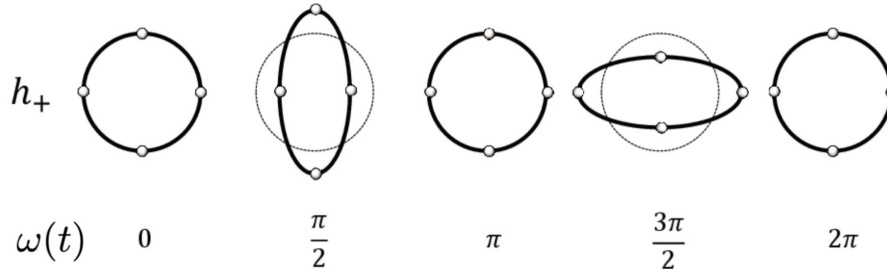
Gravitational Wave Detection

GRAVITATIONAL PERTURBATION ON MATTER

Integrating for the plus and cross polarisation leads to two set of equations:

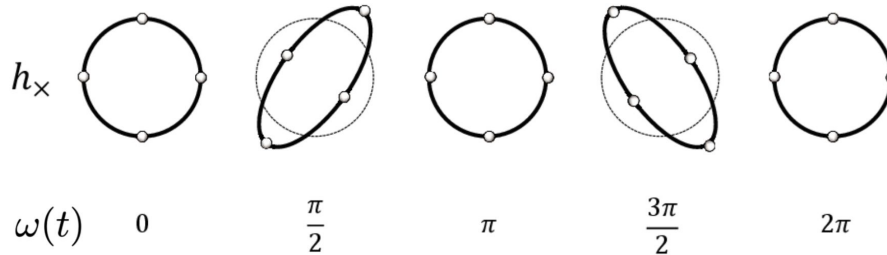
$$\delta x(t) = \frac{h_+}{2} x_0 \sin \omega t$$

$$\delta y(t) = -\frac{h_+}{2} y_0 \sin \omega t$$



$$\delta x(t) = \frac{h_\times}{2} y_0 \sin \omega t$$

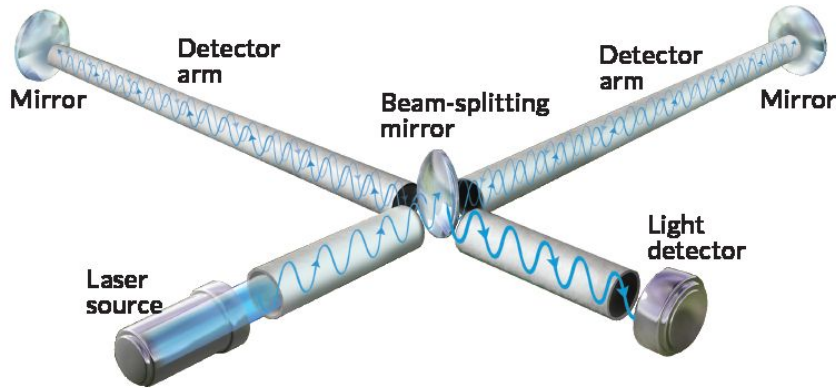
$$\delta y(t) = \frac{h_\times}{2} x_0 \sin \omega t$$



Gravitational Wave Detection

INTERFEROMETRIC DETECTORS

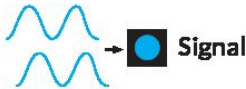
- Laser beam is sent through a beam splitter and ends up in two resonant cavities, in which luminosity is built up.
- The light is allowed to interfere where the arms join together.



Normal situation



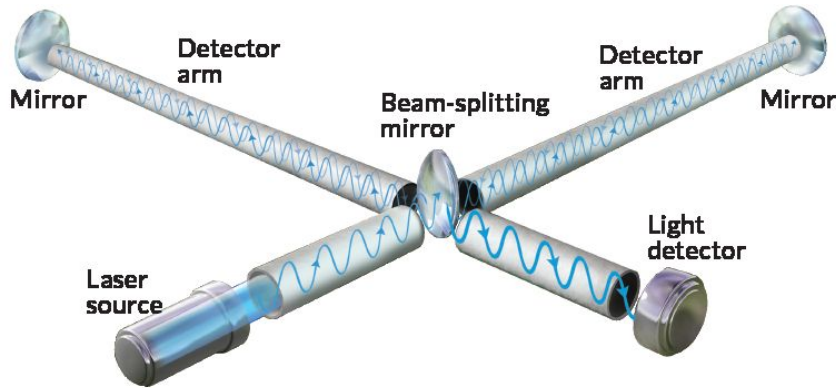
Gravitational wave detection



Gravitational Wave Detection

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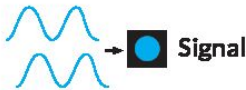
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Normal situation



Gravitational wave detection



No GW



Interference is destructive, no light hits the photodetector at the output.

GW



Light path periodically shorten in one direction and lengthen it in the other.

Gravitational Wave Detection

INTERACTION WITH TEST MASSES

The distance of a mirror with respect to the beam splitter is described by the geodesic deviation equation:

$$\ddot{L}^i = \frac{1}{2} \ddot{h}_{ij}^{TT} L^j$$



The proper distance along the x axis expands and shrinks periodically, with a fractional length change

$$\frac{\delta L}{L_c} \simeq \frac{1}{2} h_{xx}^{TT}(t, z = 0)$$

Gravitational Wave Detection

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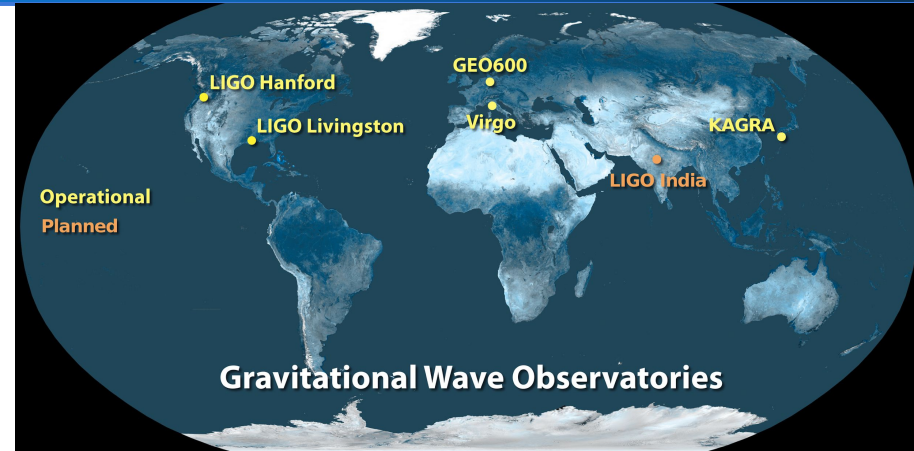
Since $h(t)$ is linear in h_+ and h_\times , it can be written through the *beam pattern functions*

$$h(t) = F_+ h_+(t) + F_\times h_\times(t)$$

Gravitational Wave Detection

Network of detector is necessary to

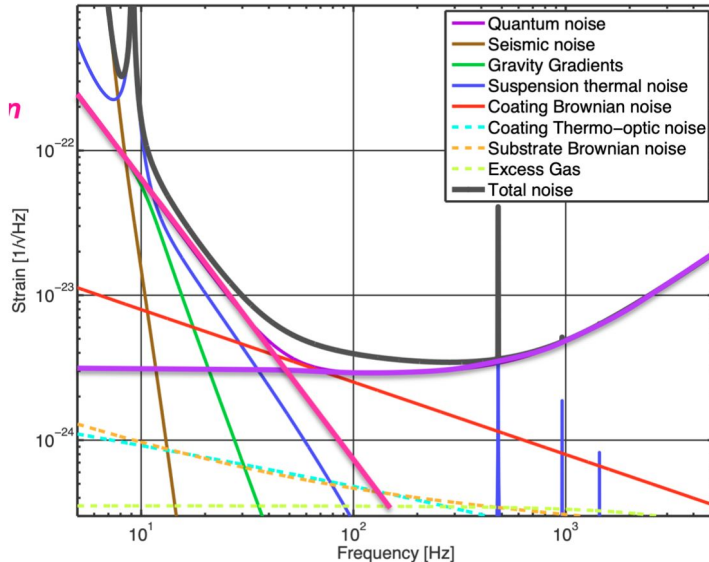
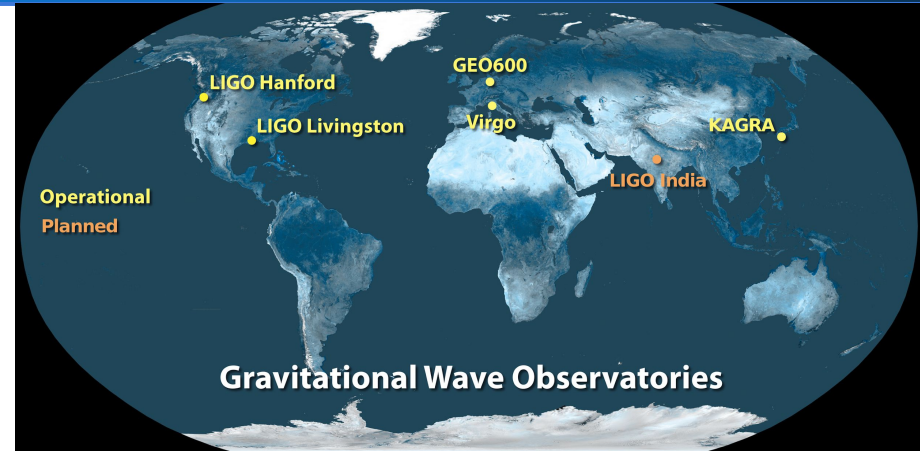
- provide directional information about the source
- separate a signal from noise confidently



Gravitational Wave Detection

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If there is a signal: measured strain is a sum of noise and signal:

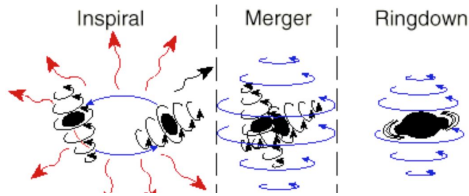
$$s(t) = n(t) + h(t)$$

instrumental noise arising from naturally occurring random processes

Gravitational Wave Sources

INSPIRALLING BINARY SYSTEMS:

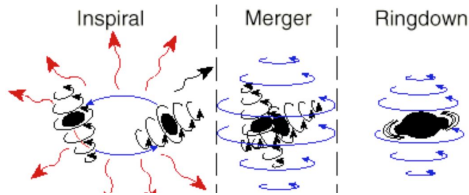
- Compact objects: neutron stars and/or black holes.
- Three stages evolution that leads to coalescence.
- Inspiral needs to be in its final stages in order for the GWs to be detectable by Earth-based interferometers.



Gravitational Wave Sources

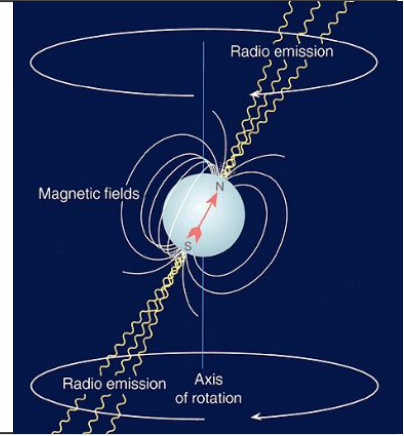
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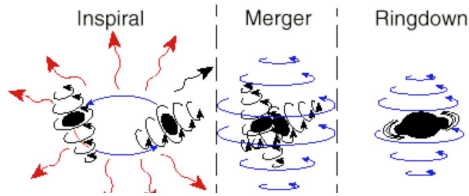
- Continuous sources of GWs such as pulsars with a non-trivial quadrupole moment.
- Evolution time is longer, GW emission weaker
- GW signal almost constant in amplitude and frequency



Gravitational Wave Sources

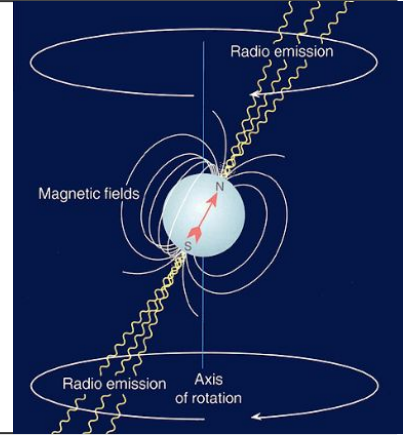
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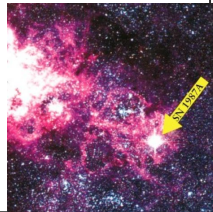
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UNMODELED BURST

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- Potential progenitors: systems such as supernovae or gamma ray bursts

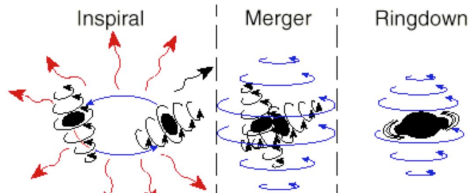
- Little is known about the details of these systems to anticipate an accurate waveform



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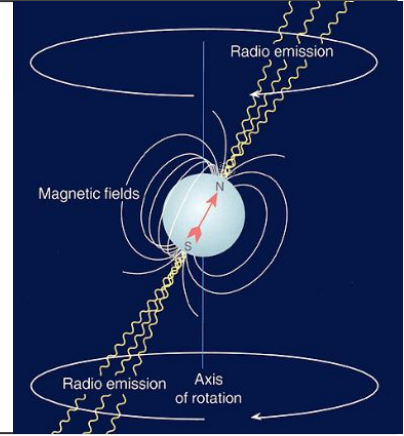
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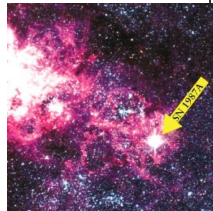
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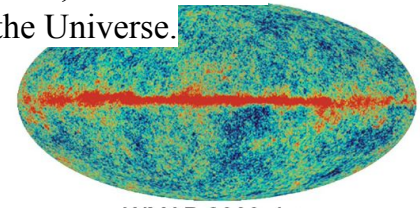
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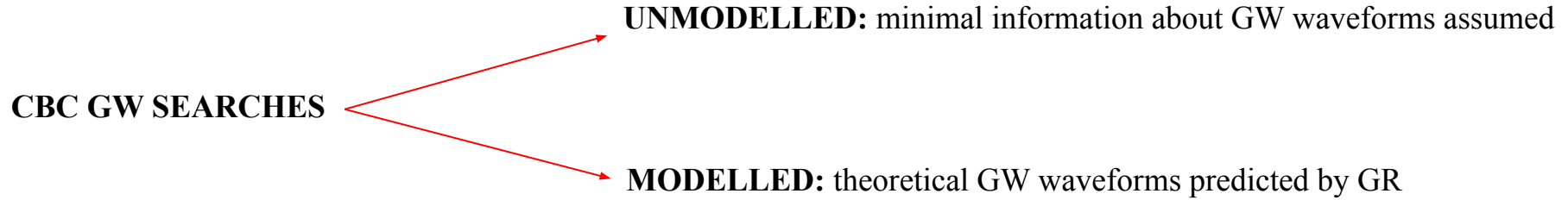


STOCHASTIC GWs

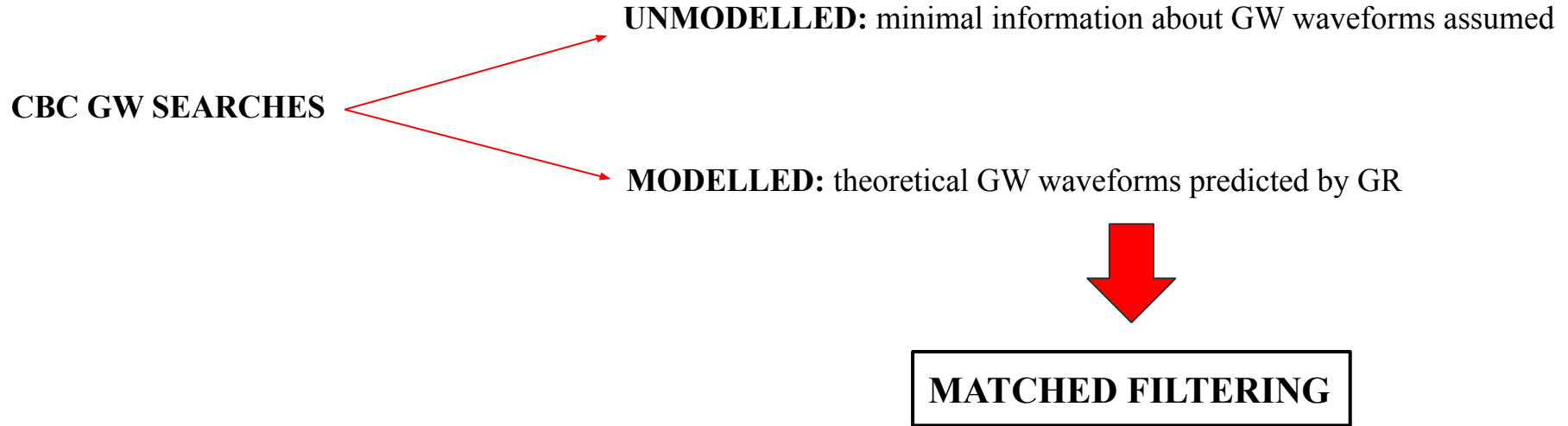
- Remnant GWs from the early evolution of the universe.
- Superposition of GWs produced by many simultaneous inspirals, bursts, or continuous signals from throughout the Universe.
- cosmic GW background



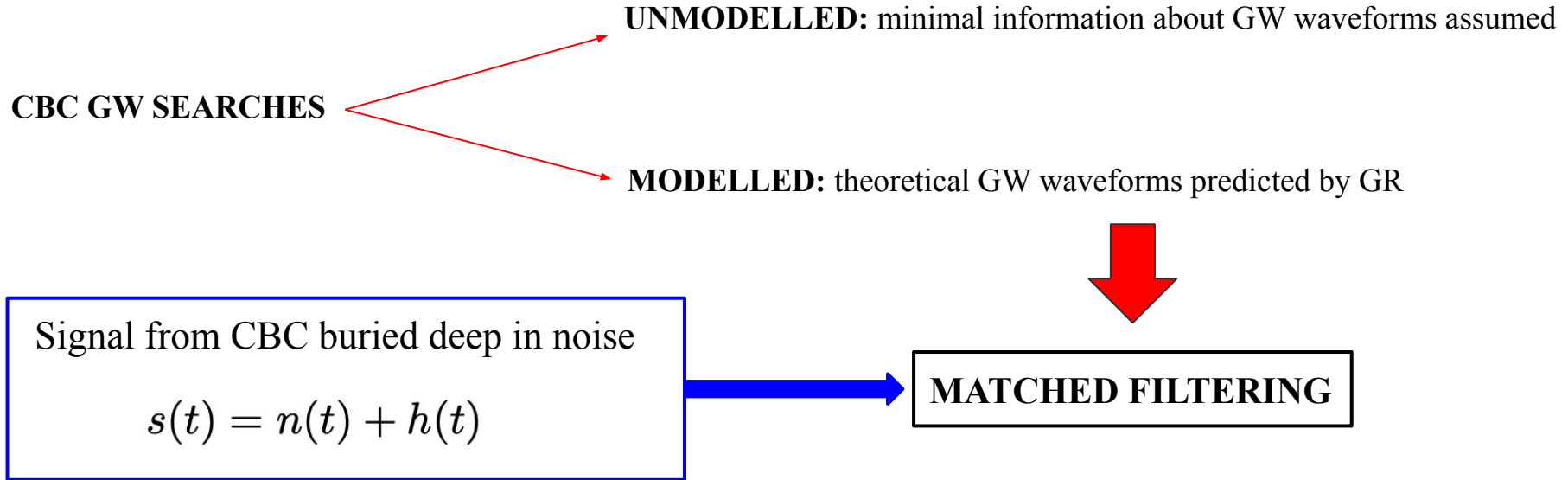
Gravitational Wave Data Analysis



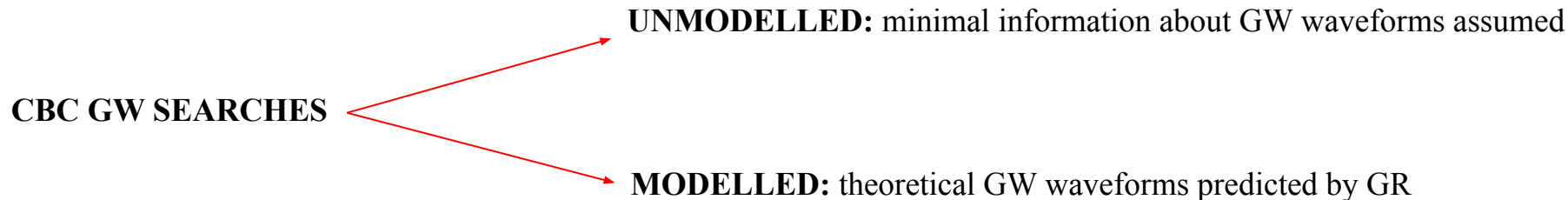
Gravitational Wave Data Analysis



Gravitational Wave Data Analysis



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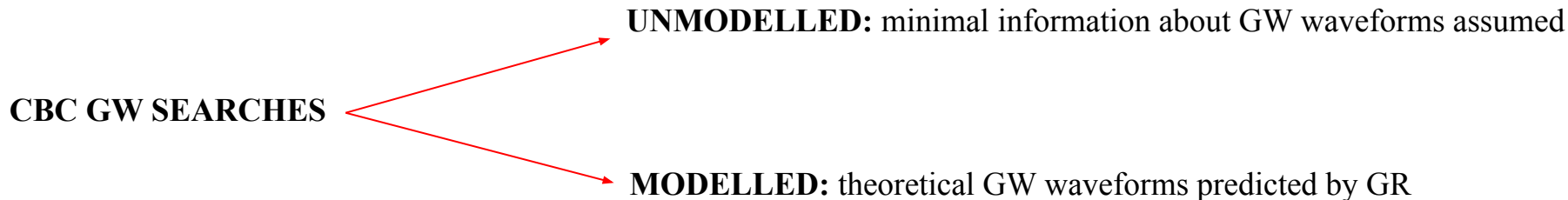
Signal from CBC buried deep in noise

$$s(t) = n(t) + h(t)$$

MATCHED FILTERING

$$(s|h)(t) = 4\text{Re} \int_{f^{low}}^{f^{high}} \frac{\tilde{s}(f) \tilde{h}_{\text{template}}^*(f)}{S_n(f)} e^{2\pi i f t} df ,$$

Gravitational Wave Data Analysis



Signal from CBC buried deep in noise

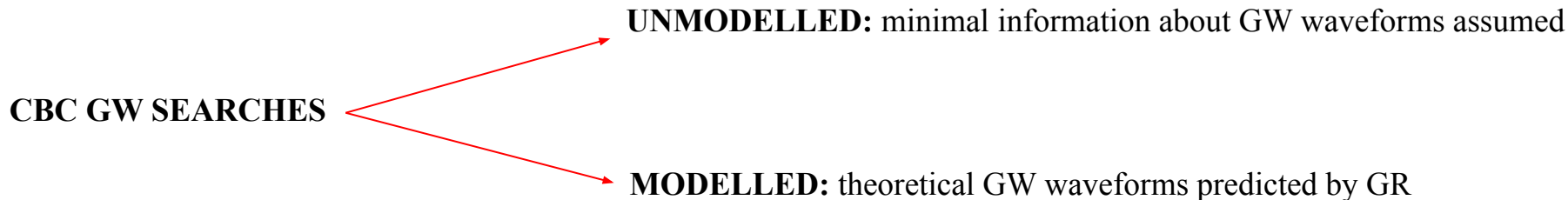
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optimal filter

Gravitational Wave Data Analysis



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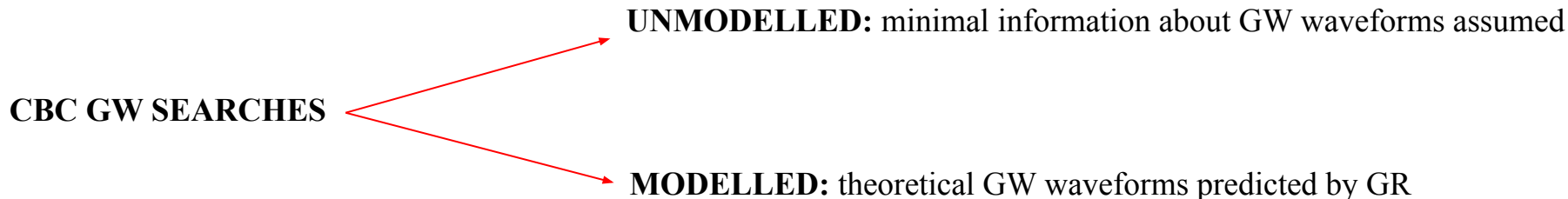
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PSD

Gravitational Wave Data Analysis



Signal from CBC buried deep in noise

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MATCHED FILTERING

**FILTER THAT MAXIMISES
THE OUTPUT OPTIMALLY
EXTRACTS THE SIGNAL
FROM NOISE!!**

$$(s|h)(t) = 4\text{Re} \int_{f_{\text{low}}}^{f_{\text{high}}} \frac{\tilde{s}(f) \tilde{h}_{\text{template}}^*(f)}{S_n(f)} e^{2\pi i f t} df ,$$

Gravitational Wave Data Analysis

Signal candidate or not?



Thresholding statistic necessary: SNR

$$\rho(t) = \frac{|(s|h)|}{\sigma_h}$$

Gravitational Wave Data Analysis

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observed filter output

Gravitational Wave Data Analysis

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root-mean-square fluctuations
of the optimal filter

Gravitational Wave Data Analysis

Signal candidate or not?



Thresholding statistic necessary: SNR

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- Proportional to the amplitude of the signal buried in the noise.
- Threshold needs to be accurately chosen to avoid high FAR while detecting signals

Gravitational Wave Data Analysis

- Parameters of astrophysical systems will not be known a-priori.
- We know approximate *waveforms* of various sources (BBH, BNS and NSBH).
- To cover the full parameter space is computationally prohibitive.
 - maximise over an overall amplitude and phase of the signal.
 - use inverse Fourier transform to evaluate the statistic as a function of time.
 - **Only intrinsic parameters - masses and spins - remain.**



Match-filter data against a predetermined collection of waveform models:
the template bank.

Gravitational Wave Data Analysis

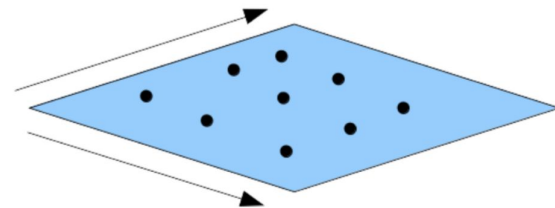
TEMPLATE BANKS

Waveforms manifold: continuous space in the component masses and spins.

Lay out template banks over parameter space: points $\lambda = \{\lambda_{intr}, \lambda_{extr}\}$ waveforms $h(f; \lambda)$

NORMALIZED
OVERLAP

$$(\hat{h}_i | \hat{h}_j) = \frac{(h_i | h_j)}{\sqrt{(h_i | h_i)(h_j | h_j)}}.$$



Gravitational Wave Data Analysis

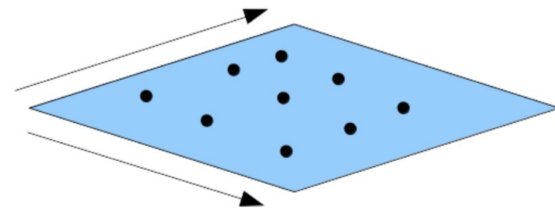
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MATCH BETWEEN
NEARBY TEMPLATES

$$\mathcal{M}(\lambda, \Delta\lambda) \equiv \max_{\Delta\lambda_{extr}} (\hat{h}(f; \lambda), \hat{h}(f; \lambda + \Delta\lambda))$$

Gravitational Wave Data Analysis

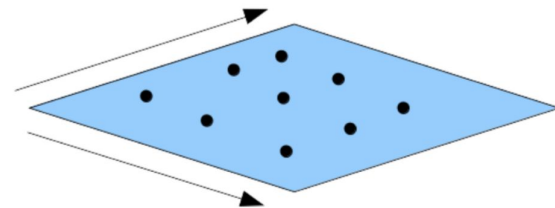
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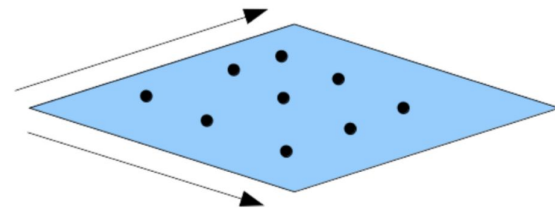
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**METRIC
WAVEFORM
MANIFOLD**

EXPAND

$$\mathcal{M}(\lambda, \Delta\lambda) \simeq 1 - g_{ij} \Delta\lambda^i \Delta\lambda^j,$$

$$g_{ij} \equiv -\frac{1}{2} \frac{\partial^2 \mathcal{M}}{\partial \Delta\lambda^i \partial \Delta\lambda^j} \Big|_{\Delta\lambda=0}$$

Gravitational Wave Data Analysis

TEMPLATE BANKS

Templates placement in parameter space such that metric distance is never larger than a preset mismatch:

$$1 - \mathcal{M} = g_{ij} \Delta \lambda^i \Delta \lambda^j$$

Choice of the set of waveform models based on *Fitting Factor*:

$$FF(h_*) = \max_{\Lambda} \mathcal{M}(\tilde{h}_*, \tilde{h}(\Lambda)) .$$

Template discrete placement \longrightarrow Fractional loss of SNR in capturing the signal with the template bank.

$$\boxed{1 - FF(h_*)}$$

Gravitational Wave Data Analysis

TEMPLATE BANKS

Effectualness: template bank is built such that no putative signal anywhere in the parameter space has

$FF < \text{the minimal match}$

- Reasonable FF value for effectual bank~97%
- FF distribution as indicator of parameter space regions with poor sensitivity.
- Systems whose GW is weaker corresponds to lowest FF:

Gravitational Wave Data Analysis

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$$FF_{\text{eff}} = \alpha^{1/3}$$

- ★ Intrinsically bright sources are favoured.
- ★ Poorly aligned and intrinsically low GW luminosity disfavoured.

$$\alpha = \frac{\int d\vec{x} F F^3(\vec{x}) \sigma^3(\vec{x})}{\int d\vec{x} \sigma^3(\vec{x})}$$

Gravitational Wave Data Analysis

CANDIDATE RANKING STATISTIC

Detector calibrated strain data contains both Gaussian and non-Gaussian noise.



Background under control

- **Data quality investigations and vetoes**
- **Consistency checks**
 - Mitigate the effect of noise transients.
 - **Chi-square test:** is signal distribution across frequency band as expected?
 - Assign an accurate statistical significance to candidate signals.

Gravitational Wave Data Analysis

CANDIDATE RANKING STATISTIC

- **Chi-square test**

Gravitational Wave Data Analysis

CANDIDATE RANKING STATISTIC

Background estimation

- Estimate search background from the data themselves.
- Time-shift technique: repeat analysis with time offset applied between detectors.
 - Background noise sample: coincidences resulting in this time-shifted data.
 - Each coincident trigger is assigned a FAR.
- Divide search spaces into different classes to deal with non-uniform background.
 - Candidate events is measured against the background from the same class.
- Hierarchical removal of confirmed signals from background to assess significance of other events.

Gravitational Wave Data Analysis

DETECTION STRATEGIES

COINCIDENT ANALYSIS

- Data from individual detectors are processed separately and then combined.
- Signals must be seen in both detectors within an allowed time window.
- Matched filtering done once per detector per template.
- Non-coincident triggers immediately rejected.
- Found coincidences are ranked by

$$\hat{\rho}_{\text{coinc}} = \sqrt{\hat{\rho}_1^2 + \hat{\rho}_2^2}.$$

- Background estimation analysis repeated, FAR computed as function of the detection statistic.

Gravitational Wave Data Analysis

DETECTION STRATEGIES

COINCIDENT ANALYSIS

Detection is claimed

Gravitational Wave Data Analysis

DETECTION STRATEGIES

COHERENT ANALYSIS

- Data from the individual detectors are combined coherently in phase and then processed.
- Computationally expensive for modelled searches.
 - Externally triggered searches: time and sky-location are known, computational cost is decreased.
- Single statistic for the full network as to construct just one more sensitive detector.
 - New ranking statistic.

Gravitational Wave Data Analysis

DETECTION STRATEGIES

COHERENT ANALYSIS

The multi-detector
log-likelihood

$$\ln \lambda = (\mathbf{s}|\mathbf{h}) - \frac{1}{2}(\mathbf{h}|\mathbf{h})$$

- Search a 7D parameter space of signals for binaries on circular orbits with and aligned spin components.
 - Maximising ~~log-likelihood over the values of the 4 waveform amplitudes~~ leads to:

$$\rho_{\text{coh}}^2 \equiv 2 \ln \lambda|_{\text{MAX}} \equiv (\mathbf{s}|\mathbf{h}_\mu) \mathcal{M}^{\mu\nu} (\mathbf{s}|\mathbf{h}_\nu)$$

- More sensitive search: more effective than the coincident search in removing background noise.

Gravitational Wave Data Analysis

DETECTION STRATEGIES

COHERENT ANALYSIS

- Sensitivity increases with the number of the detectors in the network.
- Other improvements:
 - Ranking statistic.
 - Noise reduction and rejection mechanism.

Null stream consistency: stream that contains transients which do not contribute power to the two-polarisation signal space.

$$\rho_{\text{null}}^2 = \rho_{\text{coinc}}^2 - \rho_{\text{coh}}^2$$

Gravitational Wave Data Analysis

DETECTION STRATEGIES

COHERENT ANALYSIS

- **Chi-square test**

Gravitational Wave Data Analysis

The PyGRB pipeline