

GALEX scan-mode Data

1. Description of Data

The Galaxy Evolution Explorer (*GALEX*, ?) is a space-based telescope investigating the causes and evolution of star formation in galaxies in UV band.

After the original *GALEX* was finished, the mission was extend to the *GALEX CAUSE* phase, in which the telescope was operated in a mode know as *scan-mode*. In contrast to the traditional boresight dither, the telescope traverses around the galactic plane rapidly. The *scan-mode* enables the team to complete the All-Sky UV Survey in a short time and avoid saturation caused by the high stellar density in the survey region.

Data used in this project are the raw photon data(-raw6.fits), the refined attitude solution (-asprta.fits), and the spacecraft state file (-scst.fits). This project also heavily depend on the photon list produced by *gPhoton*(<https://github.com/cmillion/gPhoton>), in which each photon event is tagged with position and time.

2. Calibration

2.1. Freedom of the telescope

2.2. Cross-correlation of photons and star catalog

To characterize the offsets of the pointing and the rotation correction of the detector, the cross-correlation of photons and star catalog are calculated.

3. Sky map construction

3.1. Count map

To construct the sky map image, the photon list (list of photons tagged with position and time) should be binned into pixels. In this project, the pixel size is 1.5 arcsec. The count map is defined as the binned photon image as shown in the upper panel of Fig. ??

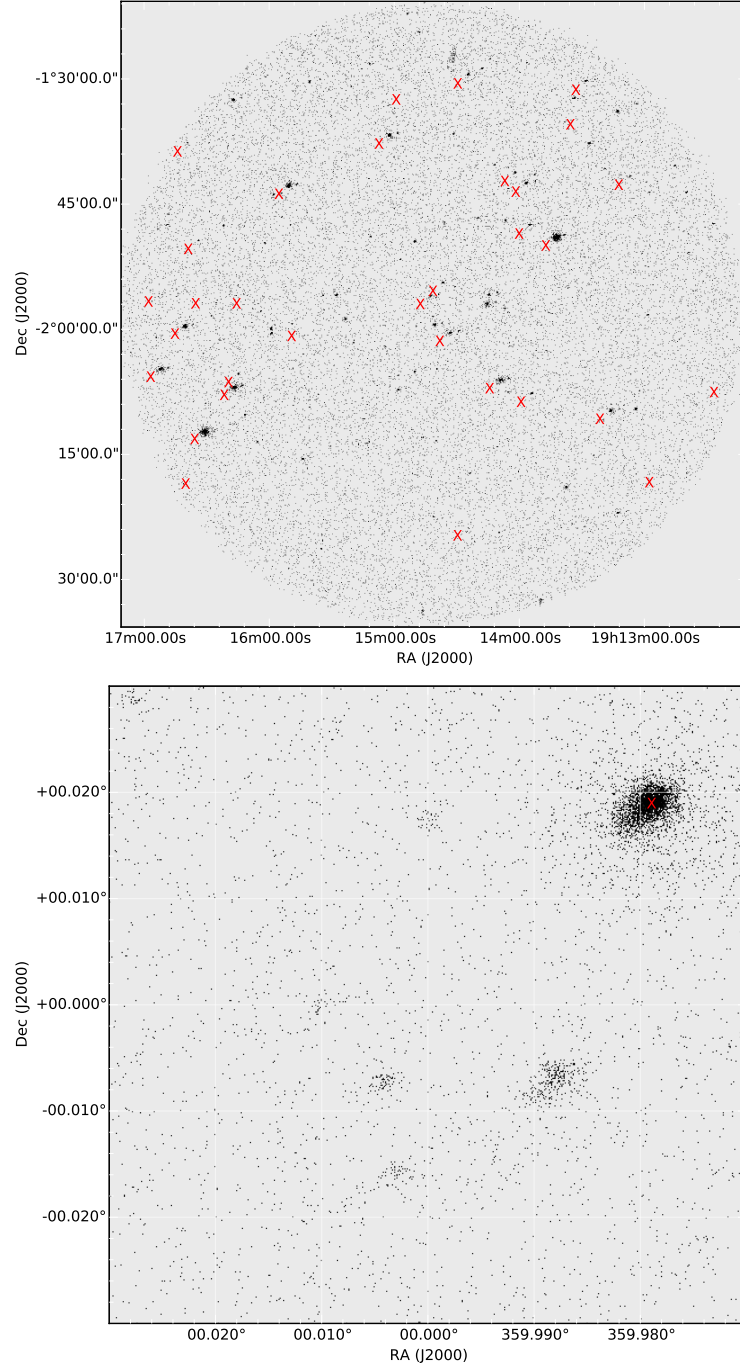


Fig. 1.— Cross-correlation of photons and star catalog.

3.2. Exposure map

The exposure map is defined as the following integral:

$$exposure = \int_{t_0}^{t_1} f(1 - D)dt \quad (1)$$

where f is the flat-field, D is the dead time recorded in the spacecraft state file (-scst.fits). The lower and upper limit of the integral t_0 , t_1 are the beginning and ending time of the scan.

3.3. Intensity map

After the count map and exposure map are constructed, the intensity map is defined as the ratio these two maps:

$$intensity = \frac{count}{exposure} \quad (2)$$

4. Source Extraction

Once the images were constructed, we used SExtractor (?) to detect sources and measure their positions and fluxes (using the FLUX_AUTO parameter). We used a modified version of the default input SExtractor configuration with the goal of minimizing false detections at the edges of the images. We found that the background subtraction step in SExtractor introduced an artificial level of noise from the image edges. To avoid this issue altogether we cut the image into a subsection that removed the outer scan edges and then ran SExtractor. We then converted the fluxes to GALEX NUV AB magnitudes using $NUV = -2.5 * \log_{10}(FLUX) + 20.08$.

The scan data had overlaps of about .25 degrees on each side, producing duplicates in the SExtractor output. To deal with this issue we looked at the overlapping regions and compared their coordinates between scans and flagged any matches. Finally, we only used SExtractor data from the sources that had the highest signal to noise.

To test the accuracy of our NUV measurements, we compared our measurements with the GALEX All Sky Imaging Survey which included regions in the Galactic plane. We gathered 16,851,560 objects within the same limits as the GALEX Plane Survey (360deg x 20deg from $abs(gb) < 10$). The magnitudes agree very well from $14 < m_{NUV} < 18$.

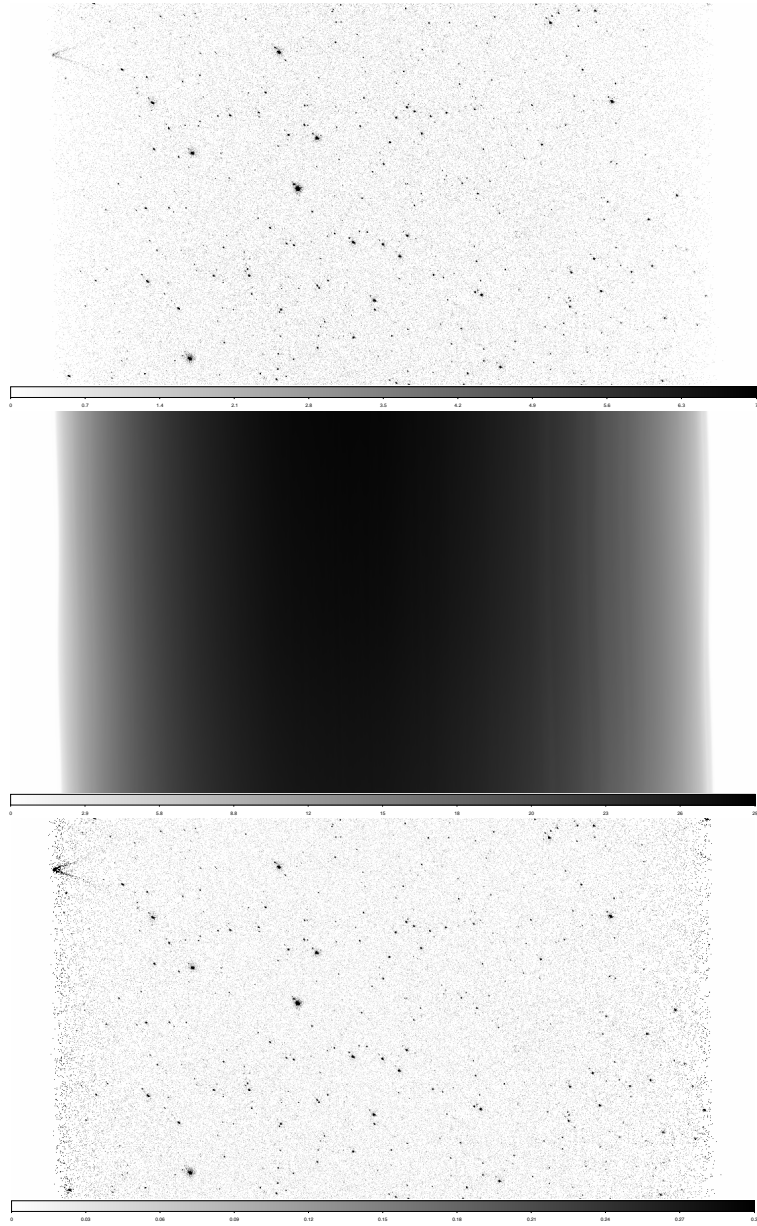


Fig. 2.—

DETECT_TYPE	CCD
DETECT_IMAGE	SAME
FLAG_IMAGE	flag.fits
DETECT_MINAREA	8
DETECT_THRESH	2.5
ANALYSIS_THRESH	2
FILTER	Y
FILTER_NAME	default.conv
DEBLEND_NTHRESH	32
DEBLEND_MINCONT	0.005
CLEAN	Y
CLEAN_PARAM	1.0
BLANK	Y
PHOT_APERTURES	10
PHOT_AUTOPARAMS	2.5, 3.5
SATUR_LEVEL	50000.0
MAG_ZEROPOINT	0.0
MAG_GAMMA	4.0
GAIN	100.0
PIXEL_SCALE	7.2
BACK_SIZE	64
BACK_FILTERSIZE	3
BACKPHOTO_TYPE	GLOBAL
BACKPHOTO_THICK	24
CHECKIMAGE_TYPE	BACKGROUND
MEMORY_OBJSTACK	2000
MEMORY_PIXSTACK	100000
MEMORY_BUFSIZE	512
SCAN_ISOAPRATIO	0.6
VERBOSE_TYPE	NORMAL