

AN ENTROPY THRESHOLD FOR STRONG $H\alpha$ AND RADIO EMISSION IN THE CORES OF GALAXY CLUSTERS

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ABSTRACT

Our *Chandra* archival study of intracluster entropy in a sample of 222 galaxy clusters shows that $H\alpha$ and radio emission from the brightest cluster galaxy are much more pronounced when the cluster’s core gas entropy is $\lesssim 30 \text{ keV cm}^2$. The prevalence of $H\alpha$ emission below this threshold indicates that it marks a dichotomy between clusters that can harbor multiphase gas and star formation in their cores and those that cannot. The fact that strong central radio emission also appears below this boundary suggests that AGN feedback turns on when the intracluster medium starts to condense, strengthening the case for AGN feedback as the mechanism that limits star formation in the Universe’s most luminous galaxies.

Subject headings: galaxies: clusters: general – X-rays: active – galaxies: conduction – galaxies: cooling flows – galaxies

1. INTRODUCTION

In recent years the “cooling flow problem” has been the focus of intense scrutiny as the solutions have broad impact on our theories of galaxy formation (see Peterson & Fabian (2006) for a review). Current models predict that the most massive galaxies in the Universe – brightest cluster galaxies (BCGs) – should be bluer and more massive than observations find, unless AGN feedback intervenes to stop late-time star formation (Bower et al. 2006; Croton et al. 2006; Saro et al. 2006). X-ray observations of galaxy clusters have given this hypothesis considerable traction. From the properties of X-ray cavities in the intracluster medium (ICM), Bîrzan et al. (2004) concluded that AGN feedback provides the necessary energy to retard cooling in the cores of clusters (see McNamara & Nulsen 2007 for a review). This result suggests that, under the right conditions, AGN are capable of quenching star formation by heating the surrounding ICM.

If AGN feedback is indeed responsible for regulating star formation in cluster cores, then the radio and star-forming properties of galaxy clusters should be related to the distribution of ICM specific entropy¹. In previous observational work (see Donahue et al. 2005, 2006; Cavagnolo et al. 2008a), we have focused on ICM entropy as a means for understanding the cooling and heating processes in clusters because it is a more fundamental property of the ICM than temperature or density alone (Voit et al. 2002; Voit 2005). ICM temperature mainly reflects the depth and shape of the dark matter potential well, while entropy depends more directly on the history of heating and cooling within the cluster and determines the density distribution of gas within that potential.

We have therefore undertaken a large *Chandra* archival project to study how the entropy structure of clusters correlates with other cluster properties. Cavagnolo et al. (2008a) presents the radial entropy profiles we have measured for a sample of 222 clusters taken from the *Chandra* Data Archive. We have named this project the Archive of Chandra Cluster Entropy Profile Tables, or *ACCEPT* for short. To characterize the ICM entropy distributions of the clusters, we fit the equation $K(r) = K_0 + K_{100}(r/100 \text{ kpc})^\alpha$ to each entropy profile. In

this equation, K_{100} is the normalization of the power-law component at 100 kpc and we refer to K_0 as the central entropy. Bear in mind, however, that K_0 is not necessarily the minimum core entropy or the entropy at $r = 0$, nor is it the gas entropy which would be measured immediately around the AGN or in a BCG X-ray corona. Instead, K_0 represents the typical excess of core entropy above the best fitting power-law found at larger radii. Cavagnolo et al. (2008a) shows that K_0 is non-zero for almost all clusters in our sample.

In this letter we present the results of exploring the relationship between the expected by-products of cooling, *e.g.* $H\alpha$ emission, star formation, and AGN activity, and the K_0 values of clusters in our survey. To determine the activity level of feedback in cluster cores, we selected two readily available observables: $H\alpha$ and radio emission. We have found that there is a critical entropy level below which $H\alpha$ and radio emission are often present, while above this threshold these emission sources are much fainter and in most cases undetected. Our results suggest that the formation of thermal instabilities in the ICM and initiation of processes such as star formation and AGN activity are closely connected to core entropy, and we suspect that the sharp entropy threshold we have found arises from thermal conduction (Voit et al. 2008, *ApJ*, in press).



This letter proceeds in the following manner: In §2 we cover the basics of our data analysis. The entropy- $H\alpha$ relationship is discussed in §3, while the entropy-radio relationship is discussed in §4. A brief summary is provided in §5. For this letter we have assumed a flat Λ CDM Universe with cosmogony $\Omega_M = 0.3$, $\Omega_\Lambda = 0.7$, and $H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$. All uncertainties are 90% confidence.

2. DATA ANALYSIS

This section briefly describes our data reduction and methods for producing entropy profiles. More thorough explanations are given in Donahue et al. (2006), Cavagnolo et al. (2008a), and Cavagnolo et al. (2008b).

2.1. X-ray

X-ray data was taken from publicly available observations in the *Chandra* Data Archive. Following standard CIAO re-

¹ In this paper we quantify entropy in terms of the adiabatic constant $K = kTn_e^{-2/3}$.

duction techniques², data was reprocessed using CIAO 3.4.1 and CALDB 3.4.0, resulting in point source and flare clean events files at level-2. Entropy profiles were derived from the radial ICM temperature and electron density profiles.

Radial temperature profiles were created by dividing each cluster into concentric annuli with the requirement of at least three annuli containing a minimum of 2500 counts each. Source spectra were extracted from these annuli, while corresponding background spectra were extracted from blank-sky backgrounds tailored to match each observation. Each blank-sky background was corrected to account for variation of the hard-particle background, while spatial variation of the soft-galactic background was accounted for through addition of a fixed background component during spectral fitting. Weighted responses which account for spatial variations of the CCD calibration were also created for each observation. Spectra were then fit over the energy range 0.7-7.0 keV in XSPEC 11.3.2ag (Arnaud 1996) using a single-component absorbed thermal model.

Radial electron density profiles were created using surface brightness profiles and spectroscopic information. Exposure-corrected, background-subtracted, point-source-clean surface brightness profiles were extracted from 5'' concentric annular bins over the energy range 0.7-2.0 keV. In conjunction with the spectroscopic normalization and 0.7-2.0 keV count rate, surface brightness was converted to electron density using the deprojection technique of Kriss et al. (1983). Errors were estimated using 5000 Monte Carlo realizations of the surface brightness profile.

A radial entropy profile for each cluster was then produced from the temperature and electron density profiles. The entropy profiles were fit with a simple model which is a power-law at large radii and approaches a constant value, K_0 , at small radii (see §1 for the equation). We define central entropy as K_0 from the best-fit model.

2.2. $H\alpha$

One goal of our project was to determine if ICM entropy is connected to processes like star formation. Here we do not directly measure star formation but instead use $H\alpha$, which is usually a strong indicator of ongoing star formation in galaxies (Kennicutt 1983). It is possible that some of the $H\alpha$ emission from BCGs is not produced by star formation (Begelman & Fabian 1990; Sparks et al. 2004; Ruszkowski et al. 2008; Ferland et al. 2008). Nevertheless, $H\alpha$ emission unambiguously indicates the presence of $\sim 10^4$ K gas in the cluster core and therefore the presence of a multiphase intracluster medium that could potentially form stars.

Our $H\alpha$ values have been gathered from several sources, most notably Crawford et al. (1999). Additional sources of data are M. Donahue's observations taken at Las Campanas and Palomar (reported in Cavagnolo et al. 2008a), Heckman et al. (1989), Donahue et al. (1992), Lawrence et al. (1996), Valluri & Anupama (1996), White et al. (1997), Crawford et al. (2005), and Quillen et al. (2007). We have recalculated the $H\alpha$ measurements from these sources assuming a Λ CDM cosmological model to place them on a more equal footing. However, they were made with a variety of apertures and in many cases may not reflect the full $H\alpha$ luminosity of the BCG. Thus, the exact levels of $L_{H\alpha}$ are not important for the purposes of this letter and we suggest the reader interpret the $L_{H\alpha}$ values as a binary indicator of multiphase gas: either

$H\alpha$ emission and cool gas are present or they are not.

2.3. Radio

Another goal of this work was to explore the relationship between ICM entropy and AGN activity. It has long been known that BCGs are more likely to host radio-loud AGN than other cluster galaxies (Burns et al. 1981; Valentijn & Billeweld 1983; Burns 1990). Thus, we chose to look for radio emission from the BCG of each ACCEPT cluster as a sign of AGN activity.

To make the radio measurements, we have taken advantage of the nearly all-sky flux-limited coverage of the NRAO VLA Sky Survey (NVSS, Condon et al. 1998) and Sydney University Molonglo Sky Survey (SUMSS, Bock et al. 1999; Mauch et al. 2003). NVSS is a continuum survey at 1.4 GHz of the entire sky north of $\delta = -40^\circ$, while SUMSS is a continuum survey at 843 MHz of the entire sky south of $\delta = -30^\circ$. The completeness limit of NVSS is ≈ 2.5 mJy and for SUMSS it is ≈ 10 mJy when $\delta > -50^\circ$ or ≈ 6 mJy when $\delta \leq -50^\circ$. The NVSS positional uncertainty for both right ascension and declination is $\lesssim 1''$ for sources brighter than 15 mJy, and $\approx 7''$ at the survey detection limit (Condon et al. 1998). At $z = 0.2$, these uncertainties equal distances on the sky of $\sim 3-20$ kpc. For SUMSS, the positional uncertainty is $\lesssim 2''$ for sources brighter than 20 mJy, and is always less than $10''$ (Bock et al. 1999; Mauch et al. 2003). The distance at $z = 0.2$ associated with these uncertainties is $\sim 6-30$ kpc. We calculate the radio power for each radio source using the standard relation $\nu L_\nu = 4\pi D_L^2 S_\nu f_0$ where S_ν is the 1.4 GHz or 843 MHz flux from NVSS or SUMSS, D_L is the luminosity distance, and f_0 is the central beam frequency of the observations.

Radio sources were found using two methods. The first method was to search for sources within a fixed angular distance of $20''$ around the cluster X-ray peak. The probability of randomly finding a radio source within an aperture of $20''$ is exceedingly low (< 0.004 for NVSS). Thus, in 222 total field searches, we expect to find no more than one spurious source. The second method involved searching for sources within 20 kpc of the cluster X-ray peak. At $z \approx 0.051$, $1''$ equals 1 kpc, thus for clusters at $z \gtrsim 0.05$ the 20 kpc aperture is smaller than the $20''$ aperture and the likelihood of finding a spurious source gets smaller. Both methods produce nearly the same list of radio sources with the differences being the very large, extended lobes of low-redshift radio sources such as Hydra A.

To make a spatial and morphological assessment of the radio emission's origins, *i.e.* determining if the radio emission is associated with the BCG, high angular resolution is necessary. However, NVSS and SUMSS are low-resolution surveys with FWHM at $\approx 45''$. We therefore cannot distinguish between ghost cavities/relics, extended lobes, point sources, re-accelerated regions, or if the emission is coming from a galaxy very near the BCG or a background/foreground source. We have handled this complication by visually inspecting each radio source in relation to the optical (using DSS I/II) and infrared (using 2MASS) emission of the BCG. We have used this method to establish that the radio emission is most likely coming from the BCG. When available, high resolution data from VLA FIRST³ was added to the visual inspection.

3. $H\alpha$ EMISSION AND CENTRAL ENTROPY

² <http://cxc.harvard.edu/ciao/guides/>

³ Faint Images of the Radio Sky at Twenty cm; <http://sundog.stsci.edu>

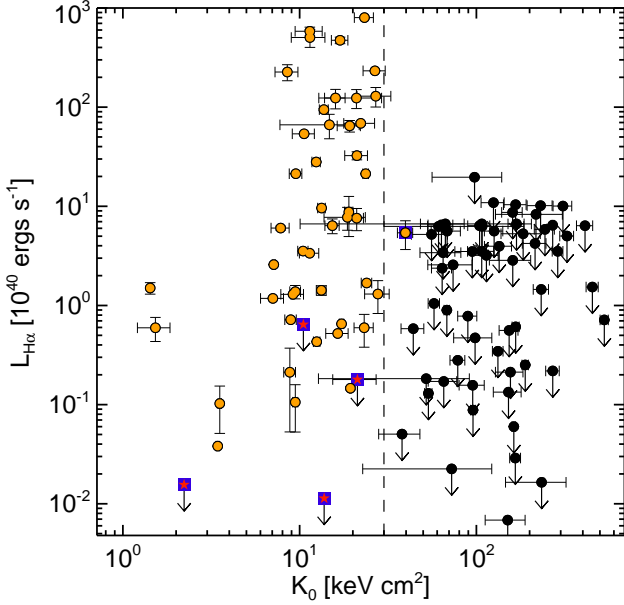


FIG. 1.— Central entropy versus $H\alpha$ luminosity. Orange circles represent detections, black circles are upper limits, and blue boxes with inset red stars or orange circles are peculiar clusters which do not adhere to the observed trend (see text). The vertical dashed line marks $K_0 = 30 \text{ keV cm}^2$. Note the presence of a sharp $H\alpha$ detection dichotomy beginning at $K_0 \lesssim 30 \text{ keV cm}^2$.

Of the 222 clusters in *ACCEPT*, we located $H\alpha$ observations from the literature for 110 clusters. Of those 110, $H\alpha$ was detected in 46, while the remaining 64 have upper limits. The mean central entropy for clusters with detections is $K_0 = 13.9 \pm 4.9 \text{ keV cm}^2$, and for clusters with only upper limits $K_0 = 128 \pm 48 \text{ keV cm}^2$.

In Figure 1 central entropy is plotted versus $H\alpha$ luminosity. One can immediately see the dichotomy between clusters with and without $H\alpha$ emission. If a cluster has a central entropy $\lesssim 30 \text{ keV cm}^2$ then $H\alpha$ emission is usually “on”, while above this threshold the emission is predominantly “off”. For brevity we refer to this threshold as K_{thresh} hereafter. The cluster above K_{thresh} which has $H\alpha$ emission (blue box with orange dot) is Zwicky 2701 ($K_0 = 39.7 \pm 3.9 \text{ keV cm}^2$). There are also clusters below K_{thresh} without $H\alpha$ emission (blue boxes with red stars): A2029, A2107, A2151⁴, EXO 0422-086, and RBS 533. These five clusters are clearly exceptions to the much larger trend. The mean and dispersion of the redshifts for clusters with and without $H\alpha$ are nearly identical, $z = 0.124 \pm 0.106$ and $z = 0.134 \pm 0.084$ respectively, and applying a redshift cut (*i.e.* $z = 0 - 0.15$ or $z = 0.15 - 0.3$) does not change the K_0 - $H\alpha$ dichotomy. Most important to note is that changes in the $H\alpha$ luminosities because of aperture effects will move points up or down in Figure 1, while mobility along the K_0 axis is minimal. Qualitatively, the correlation between low central entropy and presence of $H\alpha$ emission is very robust.

The clusters with $H\alpha$ detections are typically between $10 - 30 \text{ keV cm}^2$, have short central cooling times ($< 1 \text{ Gyr}$), and under older nomenclature would be classified as “cooling flow” clusters. It has long been known that star formation and associated $H\alpha$ nebulosity appear only in cluster cores with cooling times less than a Hubble time (Hu et al. 1985; Johnstone et al. 1987; McNamara & O’Connell 1989; Voit & Donahue 1997; Cardiel et al. 1998). However, our results suggest

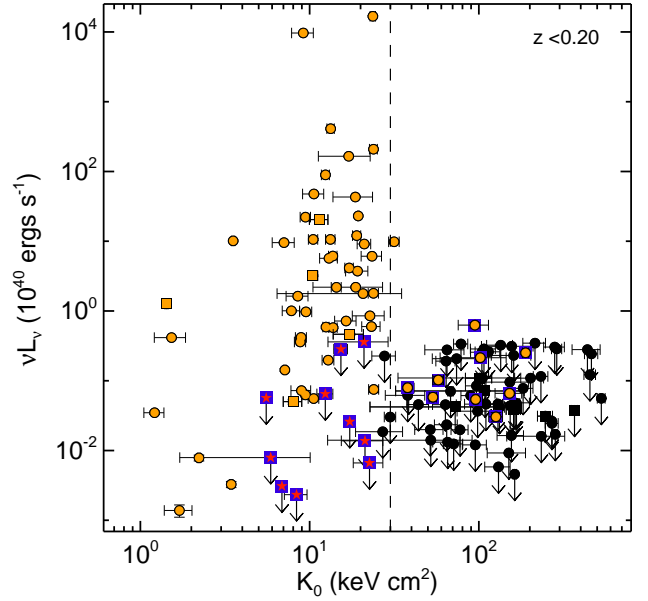


FIG. 2.— BCG radio power vs. K_0 for clusters at $z < 0.2$. Orange points represent detections, black points are non-detection upper-limits, and blue boxes with inset red stars or orange points are peculiar clusters which do not adhere to the observed trend (see text). Circles are for NVSS observations and squares are for SUMSS observations. The vertical dashed line marks $K_0 = 30 \text{ keV cm}^2$. The radio sources show the same trend as $H\alpha$: bright radio emission is preferentially “on” for $K_0 \lesssim 30 \text{ keV cm}^2$.

that the central cooling time must be at least a factor of 10 smaller than a Hubble time for these manifestations of cooling and star formation to appear. It is also very interesting that the characteristic entropy threshold for strong $H\alpha$ emission is so sharp. Voit et al. (2008) have recently proposed electron thermal conduction may be responsible for setting this threshold. This hypothesis has received further support from the theoretical work of Guo et al. (2008) showing that thermal conduction can stabilize non-cool core clusters against the formation of thermal instabilities, and that AGN feedback may be required to limit star formation when conduction is insufficient.

4. RADIO SOURCES AND CENTRAL ENTROPY

Of the 222 clusters in *ACCEPT*, 100 have radio-source detections with a mean K_0 of $23.7 \pm 9.3 \text{ keV cm}^2$, while the other 122 clusters with only upper limits have a mean K_0 of $134 \pm 52 \text{ keV cm}^2$. NVSS and SUMSS are low resolution surveys with FWHM at $\approx 45''$ which at $z = 0.2$ is $\approx 150 \text{ kpc}$. This scale is larger than the size of a typical cluster cooling region and makes it difficult to determine absolutely that the radio emission is associated with the BCG. We therefore focus only on clusters at $z < 0.2$. After the redshift cut, 135 clusters remain – 64 with radio detections (mean $K_0 = 18.8 \pm 7.6 \text{ keV cm}^2$) and 71 without (mean $K_0 = 112 \pm 45 \text{ keV cm}^2$).

In Figure 2 we have plotted radio power versus K_0 . The obvious dichotomy seen in the $H\alpha$ measures and characterized by K_{thresh} , is also present in the radio. Clusters with $\nu L_\nu \gtrsim 10^{40} \text{ erg s}^{-1}$ generally have $K_0 \lesssim K_{\text{thresh}}$. This trend was first evident in Donahue et al. (2005) and suggests that AGN activity in BCGs, while not exclusively limited to clusters with low core entropy, is much more likely to be found in clusters which have a core entropy less than K_{thresh} . That star formation and AGN activity are subject to the same entropy threshold suggests the mechanism which promotes or initiates

⁴ Not plotted because K_0 is statistically consistent with zero



one is also involved in the activation of the other. If the entropy of the hot gas in the vicinity of the AGN is correlated with K_0 , then the lack of correlation between radio power and K_0 below the 30 keV cm² threshold suggests that cold-mode accretion (Pizzolato & Soker 2005; Hardcastle et al. 2007) may be the dominant method of fueling AGN in BCGs.

We have again highlighted a few exceptions to the general trend in Figure 2: clusters below K_{thresh} without radio sources (blue boxes with red stars) and clusters above K_{thresh} with radio sources (blue boxes with orange dots). The peculiar clusters below K_{thresh} are A133, A539, A1204, A2107, A2556, AWM7, ESO 5520200, MKW4, MS J0440.5+0204, and MS J1157.3+5531. The peculiar clusters above K_{thresh} are 2PIGG J0011.5-2850, A193, A586, A2063, A2147, A2244, A3558, A4038, and RBS 461. In addition, three clusters, A2151, AS405, MS 0116.3-0115, are not plotted as their best-fit K_0 are statistically consistent with zero. Finding a few clusters in our sample without radio sources where we expect to find them is not surprising given that AGN feedback could be episodic. However, the clusters above K_{thresh} with a radio-loud source are interesting, and may be special cases of BCGs with embedded coronae. Sun et al. (2007) extensively studied coronae and found they are like “mini-cooling cores” with low temperatures and high densities. Coronae are a low-entropy environment isolated from the high-entropy ICM and may provide the conditions necessary for gas cooling to proceed. Indeed, some of these peculiar clusters show indications that a very compact ($r \lesssim 5$ kpc) X-ray source is associ-

ated with the BCG (Cavagnolo et al. 2008a).

5. SUMMARY

We have presented a comparison of ICM central entropy values and measures of BCG $H\alpha$ and radio emission for a *Chandra* archival sample of galaxy clusters. We find that below a characteristic central entropy threshold of $K_0 \approx 30$ keV cm², $H\alpha$ and bright radio emission are more likely to be detected, while above this threshold $H\alpha$ is not detected and radio emission, if detected at all, is significantly fainter. The mean K_0 for clusters with and without $H\alpha$ detections are $K_0 = 13.9 \pm 4.9$ keV cm² and $K_0 = 128 \pm 48$ keV cm², respectively. For clusters at $z < 0.2$ with BCG radio emission the mean $K_0 = 18.8 \pm 7.6$ keV cm², while for BCGs with only upper limits, the mean $K_0 = 112 \pm 45$ keV cm². While other mechanisms can produce $H\alpha$ or radio emission besides star formation and AGN, if one assumes the $H\alpha$ and radio emission are coming from these two feedback sources, then our results suggest the development of multiphase gas in cluster cores (which can fuel both star formation and AGN) is strongly coupled to ICM entropy.

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