

In my opinion you should anticipate here why you identify these clusters as systems hosting a mini-halo. In particular, state clearly already here (you mention it after) how can you exclude that the radio emission is not only related to the central AGN but also to a possible MH (e.g. based on the spectral index). To this regard, the steepness of the radio spectrum could be related to the different resolution and thus sensitivity of 74 MHz and 1.4 GHz observations. Have you taken this point into account?

## X-RAY FOLLOW-UP OF RADIO MINI-HALO CANDIDATES

**Introduction:** In a search for the most powerful steep<sup>1</sup> spectrum radio sources residing in cool core clusters, we have identified two candidate radio mini-halos (MHs), a rare class of object, in the galaxy clusters Abell 2675 (hereafter A2675;  $z = 0.071$ ) & Zwicky 808 (hereafter Z808<sup>2</sup>;  $z = 0.169$ ). The candidates are shown in Fig 1. Utilizing a sample of 750 galaxy clusters selected from the BCS/eBCS/REFLEX Surveys, just 10 objects were identified with central dominant (cD) galaxy optical line emission and an associated steep spectrum radio source with VLSS 74 MHz flux  $> 1\text{Jy}$ : 2A 0335+096, A133, A496, A2009, A2675, MKW3s, MKW8, MS 0735.6+7421, Z808, Z2701. Of the 10, A2675 and Z808 are the only objects not observed with *Chandra*, and 2A 0335+096 is the only other suspected MH. The correlation between sub-Gyr ICM core cooling time and cD line emission [1,2] suggests A2675 and Z808 are cool core clusters. The A2675 VLSS 74 MHz source is extended, has diameter  $\approx 200$  kpc,  $L_{74\text{ MHz}} \approx 10^{40}$  erg s<sup>-1</sup>,  $\alpha = -1.54$ , and surface brightness  $\mu \approx 50$   $\mu\text{Jy arcsec}^{-2}$ . Likewise, Z808's VLSS source is extended, has diameter  $\approx 450$  kpc,  $L_{74\text{ MHz}} \approx 10^{42}$  erg s<sup>-1</sup>,  $\alpha = -1.41$ , and  $\mu \approx 700$   $\mu\text{Jy arcsec}^{-2}$ . For both clusters, the 1.4 GHz emission is significantly more compact than the coincident 74 MHz emission. These are distinctly the properties of a MH in a cool core cluster. The *ROSAT* detections reveal intracluster mediums (ICMs) with  $L_X > 3 \times 10^{44}$  erg s<sup>-1</sup> and  $T_{\text{cluster}} \sim 4$  and  $\sim 7$  keV for A2675 and Z808, respectively [3]. The combination of low redshift, high  $L_X$ , and rare MH-like radio properties make A2675 & Z808 ideal targets for *Chandra*.

**Why Care About Radio Halos?** The ICM of galaxy clusters is composed of thermal and non-thermal components. The thermal component dominates clusters, and is observed via X-ray

<sup>1</sup>“Steep” spectral index defined using the VLSS & NVSS radio surveys as  $\alpha \equiv \log[S(\nu_1)/S(\nu_2)]/\log(\nu_1/\nu_2) < -1.3$ , e.g. radio sources with substantially more power at decreasing frequency.

<sup>2</sup>Z808 is catalogued as RXC J0301.6+0155.

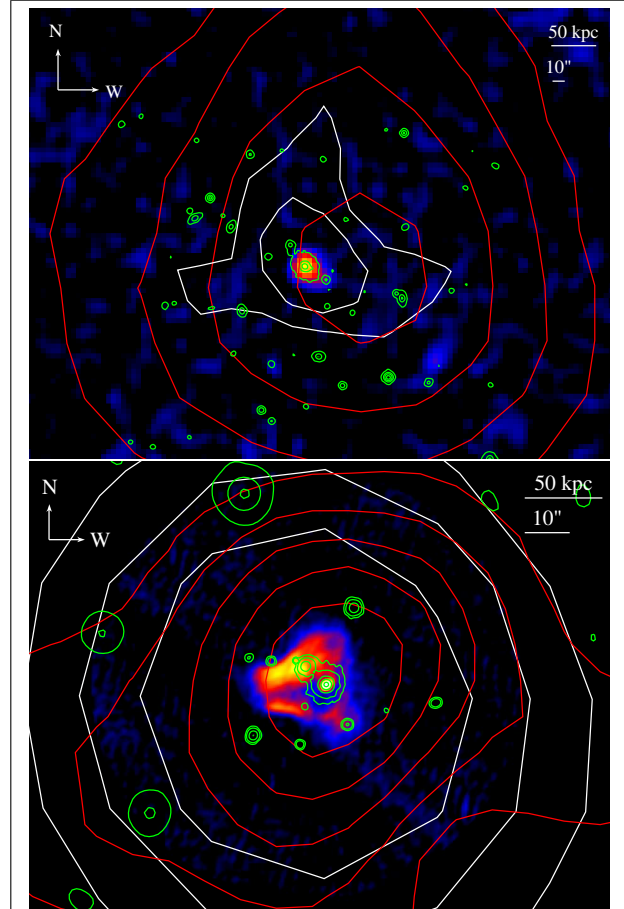


Figure 1: 1.4 GHz radio images of A2675 (top) and Z808 (bottom). Red contours are X-ray emission, green contours are optical emission, and white contours are 74 MHz radio emission.

bremsstrahlung radiation. Strong evidence exists that halo cooling of clusters with ICM cooling times  $< H_0^{-1}$  is regulated primarily by feedback from the cD active galactic nucleus (AGN) [4-6]. At the same time an AGN is influencing the evolution of the thermal component, it also contributes to the non-thermal component by injecting high-energy particles and magnetic fields ( $\vec{B}$ ) into the host environment. Evidence of this contribution comes from the relativistic jets and synchrotron emission which accompany AGN activity.

Radio MHs, like the candidates in A2675 & Z808, may be related to the process of AGN feedback. Interestingly, MHs are unilaterally found in clusters with core cooling times  $< 1$  Gyr that host a powerful ( $\gtrsim 10^{40}$  erg s<sup>-1</sup>) central radio galaxy [7,8]. Unlike high-surface brightness, low-

**General comment: do you plan more detailed radio observations of these clusters? They would be important for the scientific aims of the proposal. If yes, it should probably be mentioned.**

do you mean that rel. electrons in MHs are not simply ejected by the AGN?

what do you mean by this?

volume FR-I/FR-II sources, MHs are characterized by low radio surface brightness ( $\sim 1 - 500 \mu\text{Jy arcsec}^{-2}$  at  $< 400 \text{ MHz}$ ), large extents which fill the core ( $r \lesssim 500 \text{ kpc}$ ), and a steep radio spectrum ( $\alpha \lesssim -1$ ). **MHs ostensibly show no connection to an AGN**, and occupy a volume sufficiently large that if the radiating population had originated from a cD AGN, the particles would radiate away their energy prior to reaching the MH outskirts. This strict energetic constraint implies *in situ* particle acceleration, **possibly by diffuse  $\vec{B}$ -fields and volume-filling high-energy particles**. However, this explanation is uncertain, and even if it were not, it is unclear where the fields and particles powering MH emission originate and if, like their much larger brethren “giant” and “relic” halos, MHs are connected to mergers [9].

Given that  $\approx 50\%$  of clusters have cool cores, with  $> 90\%$  of them undergoing some level of AGN outburst or minor merger, the number of known MHs is curiously small ( $< 20$ ): 2A 0335+096, A1068, A1413, A1835, A2029, A2052, A2142, A2390, A2626, MRC 0116+111, MS 1455.0+2232, Ophiuchus, Perseus, PKS 0745-191, RX J1347.5-1145, and RX J1720.1+2638 [10-18]. We point out that all but 3 were serendipitous discoveries, that **only 6 of these meet our selection of a ROSAT detection & cD line emission, and none of those 6 have an  $\alpha < -1.1$** . Further, systematic low-frequency searches have yielded few confirmed MHs [19,20]. Four MH candidates ( $z = 0.3-0.5$ ) were recently identified with GMRT using a similar NVSS-VLSS radio-selected sample of steep spectrum sources [21]. **However, while the GMRT sources are associated with clusters, none of them are detected X-ray sources, giving  $L_X \lesssim 4 \times 10^{44} \text{ erg s}^{-1}$** . No cD emission lines are detected in three, with one **unobserved**. So while radio-selected samples exist, their utility in defining candidates for X-ray follow-up is limited.

The paucity of known MHs gives credence to the idea that they are highly transient or require very specific ICM conditions to form. While a variety of models have emerged attempting to explain MHs [e.g. 22,23], the lack of well-studied MH systems inhibits refinement of these models using observational constraints. Additionally,

if simulations incorporating AGN feedback are to reproduce the range of non-thermal sources observed in clusters, and yield insight to their importance for structure formation & evolution, then better observational constraints must be achieved. To this end, detailed X-ray analysis of MH candidate systems is vital as it enables diagnosis of cluster dynamics (e.g. with substructure like cold fronts), AGN energetics (e.g. via cavities and shocks), and how these correlate with diffuse non-thermal emission, in ways optical and radio observations are incapable [e.g. 24]. Joint high-resolution X-ray and radio study of the ICM in clusters hosting radio halos is required to better understand the link between thermal & non-thermal ICM components and to reveal the physics responsible for particle (re-)acceleration in diffuse radio halos.

ICM cold fronts (CFs) may be an especially useful tool for understanding the connection between a cool core and MH. In some MH models, ICM bulk motions & turbulence are responsible for the re-acceleration of fossil electrons which emit the diffuse synchrotron emission of a MH [e.g. 22,25]. Interestingly, turbulence and bulk motions (e.g. gas sloshing) induced by a subsonic merger event in a dense cluster core are the same processes which excite CFs [26]. A CF is detected as a constant pressure contact discontinuity where downstream gas density and temperature are higher & lower, respectively, than the upstream counterparts. That CFs appear to be long-lived, in spite of cool, high-density gas being co-spatial with gas sometimes twice as hot, indicates suppression of conduction at the CF face by  **$\vec{B}$ -fields which are likely draped over the front during its formation [27]**. Thus, the properties of a CF (i.e. size, thickness, contrast with ambient medium) provide a means for studying cool core  $\vec{B}$ -field strengths and configurations. If the  $\vec{B}$ -fields which define a CF are the same fields that produce MHs, then the study of one illuminates the other. Indeed, there are examples of MH & CFs being present in the same system with indications they are physically related [10,18,28-30]. If more examples like this could be found, a task which requires deep X-ray observations, then the cool core-CF-MH relationship could be explored

Sorry, I am veery tired... I do not see how these points strengthen and not weaken your possible MH selection. They look to me partly in contradiction with what you say before ("...These are distinct properties of a mini-halo in a cool-core cluster...")

you could also add a reference of a numerical work by Dursi & Pfrommer 2008

do you mean that no optical spectra are available?

what do you  
mean by this?

more thoroughly.

### Merger Induced MHs in A2675 & Z808?

We point-out the large companion galaxies within 15 kpc of both cDs and that the radio emission is elongated along the axis including the companions. In both clusters the optical, X-ray, and radio emission are significantly offset from each other, suggesting the cD is moving relative to the ICM. Also note the peculiar Z808 1.4 GHz morphology (sharp edge to NE of core, hole NW of core, radio plume trailing galaxy SW of cD, compressed “sandwich” appearance) and the clear alignment between the galaxies and radio features. All these features indicate recent merger activity. If so, are there CFs, shocks, or cavities in these systems which would illuminate the connection between the cool core, the MH, and previous AGN feedback? Is the MH power correlated with the cooling luminosity as is expected for models of particle re-acceleration via magnetohydrodynamic (MHD) turbulence [14]? Is there evidence of ICM turbulence in the X-ray emission (*e.g.* eddies or twisted wake like in A520)? Might there be powerful cavities in the X-ray halos of these clusters which can be used to measure the energetics of an AGN outburst and constrain how relativistic plasma is transported in the core? Spurred along by these questions, we propose X-ray observations for A2675 and Z808 with respective *Chandra* time requests of 100 ks and 80 ks. We seek to (1) determine through energetics if the radio sources have to be MHs, (2) determine if a merger has taken place, (3) probe for signs of AGN activity, (4) evaluate the connection of the steep spectrum radio source to the X-ray properties, and (5) model the A2675 & Z808 MHs using existing theories to constrain the magnetic field properties and particle content as they relate to the cool core and AGN.

**Deep X-ray Data for a Radio Halo?** The physics of how MHs are formed may be encoded in the X-ray emission of the ICM. The MH model of Gitti et al. (2004) employs MHD turbulence frozen into the gas of the cool core region to re-accelerate fossil electrons. To evaluate the Gitti model requires measurement of core properties derived from X-ray data, specifically the cool

core radius ( $r_c$ ), scaled electron gas density ( $n_c$ ), and temperature structure ( $kT_X$ ). The turbulence model of Kunz et al. (2010) also requires these parameters be known. The energy scale where MH synchrotron and inverse Compton (IC) losses are balanced by re-acceleration,  $\gamma_b$ , are related to these quantities by  $\gamma_b \propto r_c^{0.8} n_c^{-1}$ . If an AGN is responsible for the MHs, then the time for radio plasma to reach the edge of the MH should be less than the plasma synchrotron lifetime,  $t_{\text{sync}}$ . Buoyancy and sound crossing time arguments are useful in constraining the time a plasma takes to move about the ICM [2], and these calculations are based on X-ray measurements ( $t_{cs} \propto kT_X^{-1/2}$  and  $p \propto n_c kT_X$ ). Further, constraints on the ICM turbulent energy density, turbulent lengthscale, and diffusion coefficient are required, quantities which are  $n_c$  and  $kT_X$  dependent. If the cooling luminosity is much larger than the synchrotron power, then at a minimum, cooling flow powered turbulence,  $P_{\text{CF}}$ , could drive re-acceleration. Since  $P_{\text{CF}} \propto \dot{M}$ , this can only be constrained using high-SN X-ray spectral analysis of the core. We cannot neglect the possibility that merger shocks have powered the MH. In which case the energy released in the shock, determined from shock morphology and  $kT_X$  &  $n_c$  discontinuities, can be used to constrain the energy spectrum of Fermi accelerated electrons. Because of IC and synchrotron losses, the energy spectrum implies a  $t_{\text{sync}}$ , which, when set in the context of the MH morphology, determines if the halo could be powered by a shock. We will also measure radial properties of the ICM: gas mass, gravitating mass, entropy, pressure, cooling time, effective conductivity (if a strong  $T(r)$  gradient is discovered), and inferred magnetic suppression (to prevent the rapid destruction of a strong  $T(r)$  gradient).

**Request for Observations:** Our time requests are aimed at reaching temperature and density uncertainties necessary for significant detections of a typical CF, weak shock, cavities, and to collect sufficient counts for radial & spatial mapping of ICM structure. Use of ACIS-S versus ACIS-I results in 40-50% more total counts, and the ACIS-S3 FOV encloses a cluster-centric radius of

I would not mention Fermi processes and shocks here. It could be stochastic acceleration by turbulence.

1) I would move here Fig. 1 (it is easier to follow the text)  
2) In my opinion it is not clear from the green contours which are the two cDs (or probably better to say BCGs? are they really cDs with a very extended optical halo?)

This point is not really clear to me - could you give a bit more of details?

$R_{2000}$ , which is far enough out to probe for large-scale shocks or CFs. Using  $\beta$ -models fitted to the survey *ROSAT* imaging data, Cycle 12 ACIS-S count rates were determined from PIMMS. 5,000 mock surface brightness (MSB) profiles extending to  $R_{2000}$  were then generated via a Monte Carlo. Shallow baseline temperature and abundance profiles were also created,  $T(r) \propto r^{0.2}$  &  $Z(r) \propto r^{-0.2}$ , with normalizations  $T(R_{2000}) = T_{\text{cluster}}$  &  $Z(R_{2000}) = 0.3 Z_{\odot}$ . An isothermal core was also used,  $T(r < R_{7500}) = T(R_{7500})$ . For each MSB profile, cumulative counts profiles were created for exposure times ranging 40–140 ks in 20 ks steps. Bins with 2,500 counts were defined, and mean  $kT_X$  and abundance were calculated for each bin from the gradient profiles. A simulated spectrum was generated in XSPEC for each bin using an absorbed thermal model (MEKAL), the corresponding exposure time, and a normalization chosen so the spectral count rate matched the count rate predicted by the MSB profile for that bin. Source images consistent with these parameters were created in IDL, and mock *Chandra* images were created using MARX. A fixed background consisting of the PIMMS reported background and Galactic soft emission taken from RASS R12 & R45 bands was included in all analysis.

Radial profiles were extracted from the simulated data to estimate the mean uncertainties for each exposure time. We find exposure times of 100 ks for A2675 and 80 ks for Z808 return the best uncertainties per unit time while maintaining the highest signal-to-noise such that weighted Voronoi tessellation maps with a minimum binned spatial resolution of  $\sim 5\text{--}10''$  can be created. For A2675, the mean uncertainties will be  $\Delta T_X \pm 0.3\text{--}0.6$  keV and  $\Delta n_c \pm 10\%$ , and for Z808,  $\Delta T_X \pm 0.5\text{--}0.8$  keV and  $\Delta n_c \pm 12\%$ . Our simulations indicate we will be sensitive at  $> 2\sigma$  to CFs with temperature jumps  $> 1.5$  and density decreases  $> 1.3$ , which are typical values for CFs. Conversely, based on the Rankine-Hugoniot jump conditions, we should detect a  $M > 1.2$  shock at  $\gtrsim 2\sigma$ . We will also be able to measure  $\vec{B}$ -field strengths to  $\approx 3\mu\text{G}$ .

We simulated the presence of cavities in the

ICM of our targets by carving out voids in the source images and creating new mock observations with MARX. Cavity decrements, their geometries, and distances from the cluster core span a wide-range of values, thus we placed voids, obeying the radially dependent axial ratio and SB dimming relations presented in [5], at a variety of radii. For spherical plane-of-the-sky cavities, our observations should be sensitive to  $> 10\%$  decrements at  $r < 100$  kpc and  $> 30\%$  decrements at  $100 < r < 200$  with no detections expected beyond. Our mock observations loosely indicate we will be sensitive to cavity powers of  $10^{41\text{--}46}$  erg s $^{-1}$ . Cavity age estimates (important for comparison with  $t_{\text{sync}}$ ) will be uncertain to  $\pm 1$  Myr.

- (1) Crawford et al. *MNRAS*-306-857-‘99.
- (2) Cavagnolo et al. *ApJ*-683-L107-‘08.
- (3) Ebeling et al. *MNRAS*-301-881-‘98.
- (4) Croton et al. *MNRAS*-365-11-‘06.
- (5) McNamara & Nulsen. *ARA&A*-45-117-‘07.
- (6) Peterson & Fabian. *Phys. Rep.*-427-1-‘06.
- (7) Ferrari et al. *Spa. Sci. Rev.*-134-93-‘08.
- (8) Giovannini et al. *A&A*-507-1257-‘09.
- (9) Feretti & Giovannini. *LNP*-740-143-‘08.
- (10) Baum & O’Dea. *MNRAS*-250-737-‘91.
- (11) Burns et al. *ApJ*-388-L49-‘92.
- (12) Giovannini & Feretti. *NewAst*-5-335-‘00.
- (13) Bacchi et al. *A&A*-400-465-‘03.
- (14) Gitti et al. *A&A*-417-1-‘04.
- (15) Gitti et al. *A&A*-470-L25-‘07.
- (16) Govoni et al. *A&A*-499-371-‘09.
- (17) Bagchi et al. *MNRAS*-399-601-‘09.
- (18) Mazzotta & Giacintucci. *ApJ*-675-L9-‘08.
- (19) Venturi et al. *A&A*-463-937-‘07.
- (20) Venturi et al. *A&A*-484-327-‘08.
- (21) van Weeren et al. *A&A*-508-75-‘09.
- (22) Gitti et al. *A&A*-386-456-‘02.
- (23) Pfrommer & Enßlin. *A&A*-413-17-‘04.
- (24) Wong et al. *ApJ*-682-155-‘08.
- (25) Kunz et al. arXiv:1003.2719K, ‘10.
- (26) Ascasibar & Markevitch. *ApJ*-650-102-‘06.
- (27) Markevitch & Vikhlinin. *PhyRep*-443-1-‘07.
- (28) Markevitch et al. *ApJ*-541-542-‘00.
- (29) Zhao et al. *ApJ*-416-51-‘93.
- (30) Mazzotta et al. *ApJ*-596-190-‘03.

**Previous *Chandra* Programs:**

PI Cavagnolo, GO Cycle 10: “The Hyper-luminous Infrared Galaxy IRAS 09104+4109: An Extreme Brightest Cluster Galaxy.” A detailed study of IRAS 09104+4109 utilizing a 75 ks observation from Cycle 10 is presented in Cavagnolo et al. (in prep for ApJ).

PI Cavagnolo maintains the Archive of Chandra Cluster Entropy Profile Tables (ACCEPT) database and is adding 84 galaxy clusters (156 observations) to the 241 clusters currently in the database. As a result of maintaining the database, PI Cavagnolo has reduced & analyzed 752 CXO observations ( $\sim 16$  Msec of data) along with over 80,000 spectra.