

Massive Black Hole recoil in high resolution hosts

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The final inspiral and coalescence of a black hole binary can produce highly beamed gravitational wave radiation. To conserve linear momentum, the black hole remnant can recoil with “kick” velocity $v_{\text{kick}} \leq 4000$ km/s. We present two sets of full N-body simulations of recoiling massive black holes (MBH) in high-resolution, non-axisymmetric potentials. The host to the first set of simulations is the main halo of the Via Lactea I simulation (Diemand et al. 2007). The nature of the resulting orbits is investigated through a numerical model where orbits are integrated assuming an evolving, triaxial NFW potential, and dynamical friction is calculated directly from the velocity dispersion along the major axes of the main halo of Via Lactea I. By comparing the triaxial case to a spherical model, we find that the wandering time spent by the MBH is significantly increased due to the asphericity of the halo. For kicks larger than 200 km/s, the remnant MBH does not return to the inner 200 pc within 1 Gyr, a timescale an order of magnitude larger than the upper limit of the estimated QSO lifetime. The second set of simulations is run using the outcome of a high-resolution gas-rich merger (Mayer et al. 2007) as host potential. In this case, a recoil velocity of 500 km/s cannot remove the MBH from the nuclear region.

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1 Introduction

In the context of the currently favored Λ CDM cosmogony, where large halos are assembled through the hierarchical merging and accretion of small progenitors, MBH mergers should be common. While dynamical friction against the dark matter and baryons can only form binaries with separations of ~ 1 pc, other dynamical process such as stellar ejection and gas drag can efficiently reduce the binary orbital separation down to ~ 0.001 pc, the regime where gravitational wave radiation dominates the orbital energy loss. This radiation is typically anisotropic due to asymmetries associated with the black holes’ mass and spin, causing the center of mass of the system to recoil in order to balance the linear momentum carried away by the gravitational wave radiation (Bekenstein 1973; Fitchett & Detweiler 1984; Favata et al. 2004). The “kick” velocity of the remnant depends on the mass ratio ($m_2/m_1 < 1$) and the spin parameters (a_1, a_2) of the binary, but not on the total mass of the system. Early estimates of the recoil velocity, framed in the post-Newtonian regime for non-spinning, unequal mass black holes (Fitchett 1983; Redmount & Rees 1989), yielded velocities in the range $100 < v_{\text{kick}} < 500$ km/s which were successfully reproduced by numerical data (Baker et al. 2006b; González et al. 2007b; Herrmann et al. 2007). Simulations with varying (arbitrary) spin orientations (Campanelli et al. 2007b) and mass ratios (Baker et al. 2008) show

that recoil velocities can be significantly larger, reaching $v_{\text{kick}} \sim 2000$ km/s (Campanelli et al. 2007a; González et al. 2007b) and are predicted to be as large as $v_{\text{kick}} \sim 4000$ km/s for unequal mass ($m_2/m_1 = 1/3$), maximally spinning black holes (Campanelli et al. 2007a). In order to address the issue of whether off-nuclear QSO/AGN are common and detectable, we ought to understand the dynamics of their orbits. The radial orbit of a recoiling MBH in a spherically symmetric potential was studied analytically by Madau & Quataert (2004). They showed that large kicks (400 km/s) can displace MBHs a few tens of kiloparsecs away from the center of a Milky-Way size stellar bulge. After the kick, the MBH undergoes several oscillations before decaying back to the bottom of the potential. Most of the orbital energy is lost during the MBH passages through the center, where dynamical friction is most efficient. Guandris & Merritt (2008) substantiated these results by performing direct summation N-body simulations of MBH recoil in a spherically symmetric galaxy, modeled as a binary-depleted core-Sérsic profile (Graham et al. 2003). They find that after reaching the core radius, the MBH and the core experiences damped oscillations with bulge’s center of mass, which decelerate the remnant MBH until it reaches thermal equilibrium with the surrounding stars. They estimate that this final stage may take up to 1 Gyr in large galaxies. As we show in Sect. 2.2, the MBH wandering time can be largely increased if the host potential is triaxial, while the presence of gas can significantly damp the overall MBH motion.

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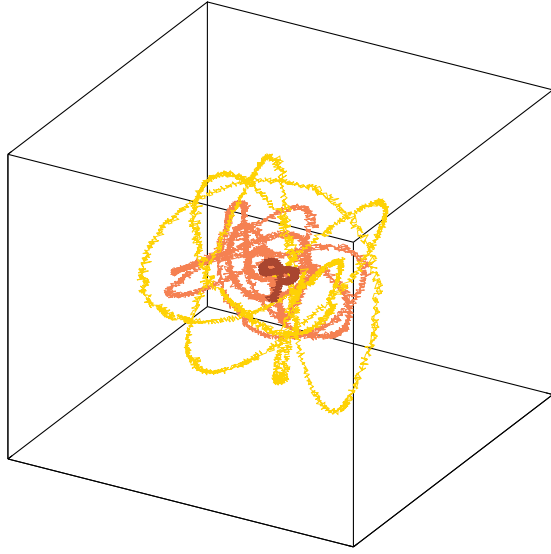


Fig. 1 (online colour at: www.an-journal.org) Resulting 3D orbit corresponding to simulation VL080. The first 0.5 Gyr are plotted in yellow, the following 0.5 Gyr in orange and the remaining 0.1 Gyr is plotted in dark red. The box has side length $L = 1.2$ kpc.

After a major merger, however, the remnant dark matter halo tends to be prolate (e.g. Novak et al. 2006), a factor that can play a major role for large kicks.

In the following we carry out N-body simulations of recoiling MBH in two high-resolution non-axisymmetric potentials: the Via Lactea I simulation's main halo, a triaxial dark-matter-only host with known triaxiality (Kuhlen et al. 2007) and evolution parameters (Diemand et al. 2007), and the gas-rich galaxy merger remnant described in Mayer et al. (2007). In the dark-matter only case, we present a numerical model that successfully reproduces the main features of the MBH orbits, and can therefore be used to estimate the fallback time and maximum displacement for any given kick velocity (Guedes et al. 2008, in preparation).

2 Recoil in a triaxial dark matter halo

2.1 Simulations

Simulations of recoil in a triaxial dark matter halo are run using the entire volume of the Via Lactea I simulation (hereafter VLI), a periodic box of size $L = 90$ Mpc. The high-resolution region contains over 230 million particles of mass of $m_p = 2.1 \times 10^4 M_\odot$. The exquisite resolution of VLI allows us to adopt the mass of SgrA* $M_\bullet = 3.7 \times 10^6 M_\odot$ (Ghez et al. 2005) for the MBH and a force softening length $\epsilon = 90$ pc. The five halo + MBH systems are evolved for 1.1 Gyr using PKDGRAV (Stadel 2001), a tree algorithm that includes up to hexa-decapole moments to reach high accuracy in the force calculation. The five runs correspond to

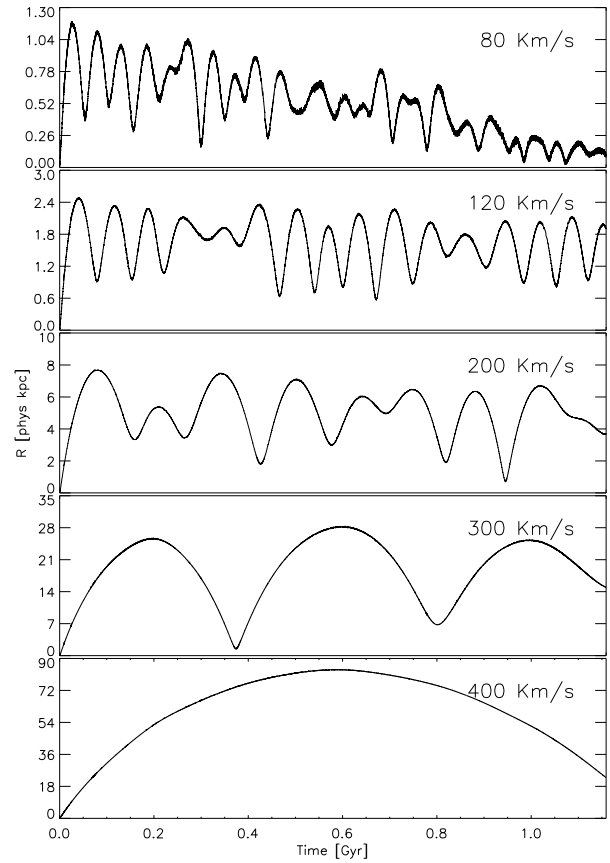


Fig. 2 Simulations VL080–VL400. Resulting orbits of recoiling MBHs from N-body simulations in the VLI main halo.

five kick velocities $v_{\text{kick}} = 80, 120, 200, 300$, and 400 km/s for runs labeled VL080 to VL400 respectively.

We place MBH at the densest point of the halo, with phase-space coordinate w_c , at an initial redshift $z_i = 1.54$, 300 Myr after the halo suffered its last major merger. At this epoch the halo is characterized by $R_{200} = 187$ kpc and $M_{200} = 1.02 \times 10^{12} M_\odot$, where R_{200} is defined as the radius within which the enclosed average density is 200 times the critical density. The MBH particle is tracked at every time-step and its position and velocity are measured with respect to w_c .

A 3D rendition of the orbit for simulation VL080 is shown in Fig. 1. Orbits for simulations VL080–VL400 are shown in Fig. 2. These orbits sample a large volume of VLI and reveal the asphericity of the main halo which causes angular momentum transfer to the y and z components of the MBH velocity. As will be shown in more detail in Guedes et al. (2008, in prep.), this increases the MBH wandering time, especially in cases where the kick is large enough to leave the baryonic component behind. For simulation VL080, which features the smallest kick velocity $v_{\text{kick}} = 80$ km/s, we are able to obtain both a maximum displacement of 1.2 kpc and a fall back time of 1.1 Gyr and therefore it can

be used to constrain the semi-analytical model described on Sect. 2.2.

2.2 Triaxial model

The orbit of a MBH can be characterized by a conservative force due to the dark matter potential $\nabla\Phi$, and a damping dynamical friction term. This generates a system of six coupled differential equations that can be separated as follows:

$$\dot{\mathbf{r}} = \mathbf{v}, \quad (1)$$

$$\dot{\mathbf{v}} = -\nabla\Phi + \mathbf{f}_{\text{DF}}.$$

The triaxial dark matter potential is modeled as modified NFW profile (Navarro, Frenk & White 1997),

$$\Phi_{\text{NFW}}^{\text{T}} = -\frac{GM_{200}}{r_s f(c)} \frac{\ln(1+x_T)}{x_T}, \quad (2)$$

$$r_T = \left(x^2 + \frac{y^2}{p^2} + \frac{z^2}{q^2}\right)^{1/2}, \quad (3)$$

where $x_T = r_T/r_s$, with r_s the scale radius, and p and q are the triaxiality parameters.

The classical Chandrasekhar dynamical friction formula is not valid in a triaxial system, because the velocity dispersion is non-isotropic and the velocity distribution deviates from Maxwellian. We adopt the Pesce et al. (1992) generalization of the dynamical friction formula for non-axisymmetric systems,

$$\mathbf{f}_{\text{DF}} = -\Gamma_1 v_1 \hat{e}_1 - \Gamma_2 v_2 \hat{e}_2 - \Gamma_3 v_3 \hat{e}_3, \quad (4)$$

where v_i is the component of the black hole velocity along the principal axis \hat{e}_i of the halo's velocity dispersion tensor, and Γ_i are the dynamical friction coefficients given by (Pesce et al. (1992); Vicari et al. (2007))

$$\Gamma_i = \frac{2\sqrt{2}\pi G^2 \rho(\mathbf{r}, z) \ln \Lambda (M_{\text{BH}} + m_p)}{\sigma_1^3} \times B_i(\mathbf{v}, \sigma), \quad (5)$$

$$B_i = \int_0^\infty \frac{\exp\left(\sum_{k=1}^3 -\frac{v_k^2/2\sigma_k^2}{\epsilon_k^2 + u}\right)}{\sqrt{(\epsilon_1^2 + u)(\epsilon_2^2 + u)(\epsilon_3^2 + u)}} \frac{1}{\epsilon_i^2 + u} du, \quad (6)$$

with $\epsilon_i = \sigma_i/\sigma_1$ and σ_1 is the largest eigenvalue.

The six components of the symmetric velocity dispersion tensor, defined as $\sigma_{ij}^2 \equiv \langle v_i v_j \rangle - \langle v_i \rangle \langle v_j \rangle$ are measured as a function of radius for every output of the VLI simulation. To get the three principal velocity dispersions, we diagonalize σ_{ij}^2 , obtaining the principal eigenvalues σ_1^2 , σ_2^2 , and σ_3^2 as a function of radius. To simplify our calculations, we perform analytical fits to σ_k^2 ,

$$\sigma_k^2(r, z) = A_k \times \frac{1 + D_k r^{m_k}}{1 + B_k r^{n_k}} e^{-r/C_k}, \quad (7)$$

where the best fit parameters A, B, C, D, n, m vary as a function of redshift.

Several studies have shown that a constant value for the Coulomb logarithm, $\ln \Lambda$, does not match simulations results. We adopt the Maoz (1993) formalism

$$\ln \Lambda \rightarrow \int_{r_{\min}}^{R_{\max}} \frac{\rho(r)}{\rho_0 r} dr, \quad (8)$$

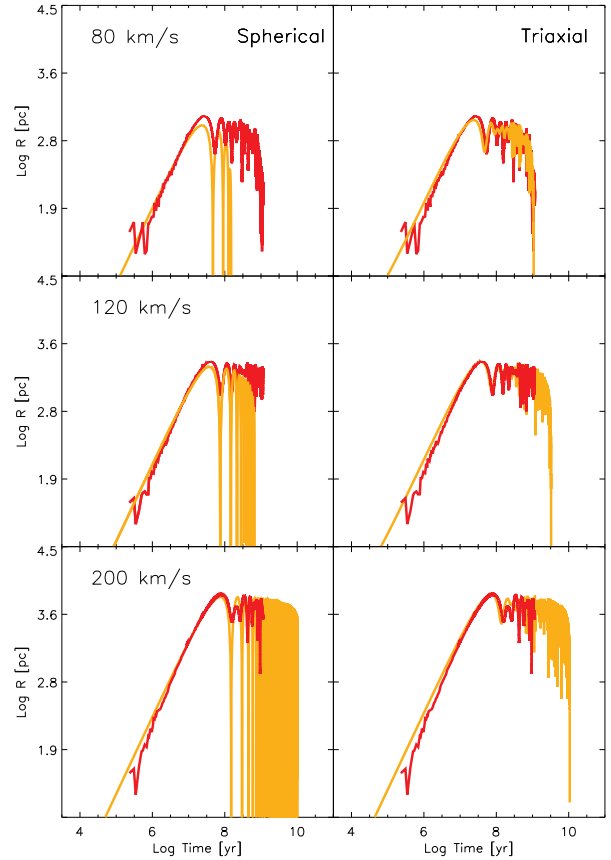


Fig. 3 (online colour at: www.an-journal.org) Comparison of spherical and triaxial models of the MBH orbits after suffering a gravitational wave recoil. The N-body simulation results (red) are superposed to the numerical model results (orange).

where ρ_0 is the interior density, $R_{\max} = r_s$, and we solve for $r_{\min} = GM_{\bullet}/\sigma^2(r)$, the sphere of influence of the MBH. The resulting orbits are shown in Fig. 3 (right). Contrary to the spherical model (left), the orbits in a triaxial halo are non-radial, which keeps the MBH orbiting away from the center of the host, reducing the mean contribution of dynamical friction to the motion and extending the wandering time. This model qualitatively reproduces the results of the N-body simulations described in Sect. 2.1.

3 Recoiling MBHs in a gas merger

We present the preliminary results of a simulation of a recoiling MBH of mass $M_{\bullet} = 3.6 \times 10^7 M_{\odot}$ in a multi-component galaxy using the SPH+N-body code GASOLINE (Wadsley, Stadel & Quinn 2004). The host is the merger remnant of a high-resolution simulation by Mayer et al. (2007), which features a gas particle mass of $m_p = 3000 M_{\odot}$, a gravitational softening $\epsilon = 2$ pc, and a central MBH in each of the progenitor galaxies. The nuclear region of the merger remnant contains 2 million particles, contains a

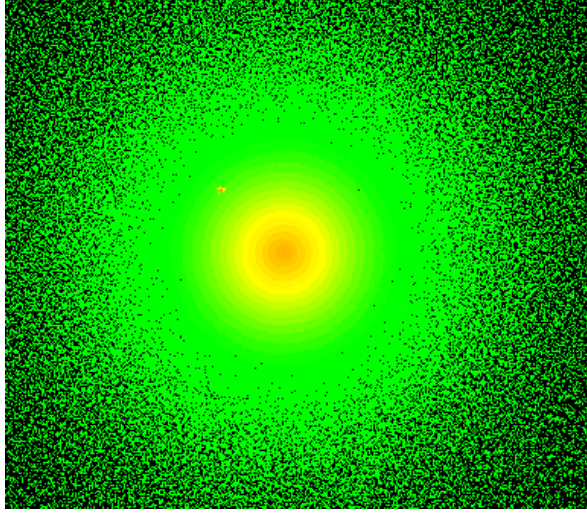


Fig. 4 (online colour at: www.an-journal.org) Projected density plot of the nuclear region of the merger remnant. A wandering MBH is seen at apocenter as an over-density near the nuclear region at $t = 5.4$ Myr after suffering a 500 km/s recoil kick. The image corresponds to an area of $30 \text{ pc} \times 30 \text{ pc}$.

mass of $M_g = 3 \times 10^9 M_\odot$ and a density that ranges from $\rho = 10^{-2}$ – 10^5 atoms/cm^3 . By the end of their simulation, the MBHs have formed a binary with semi-major axis equal to the softening length of the simulation.

We take the simulation from this point forward. We remove both MBHs and evolve the simulation until the nucleus becomes smooth, place our MBH at the center of mass of the system, and kick it with recoil velocity $v_{\text{kick}} = 500 \text{ km/s}$. In this case, the MBH was displaced only 30 pc from the center of mass. A simple Bondi accretion estimate yields an accretion rate of $\dot{M}_{\text{acc}} = 0.06 M_\odot/\text{yr}$, and therefore the MBH can become an active AGN during its wandering time.

As seen in Fig. 4 the MBH produces an over-density signature on the nuclear gas at the apocenter of its orbit, a feature that can be associated with the emission of X-ray radiation (e.g. Fujita 2008).

4 Discussion

If MBH mergers are common in the context of ΛCDM , then the “rocket effect” should also be common, especially at high-redshift where the bulk of the mass assembly occurs. Because recoiling MBHs carry gas and stars with them, the detection of off-nuclear QSOs is possible by finding velocity shifts between the narrow line region associated with gas accreting onto the MBH, and the broad line emission associated with the galaxy left behind (Bonning & Shields 2007). The best off-nuclear QSO candidate today is thought to be powered by a recoiling SMBH of mass $M_\bullet = 6 \times 10^8 M_\odot$ which travels at a velocity of 2650 km/s with respect to the host galaxy (Komossa et al. 2008).

In order to assess the detectability of so-called naked QSOs, we ought to understand the dynamics of the recoiling MBH orbits and the environment provided by the host potential. While detectability requires the presence of an accretion disk around the remnant MBH, the medium of its host galaxy can damp its motion significantly, disfavoring its detection. However, if the kick is large enough that, while keeping an accretion disk and stellar sphere on influence, the MBH can reach as far as the (generally) prolate dark matter halo, dynamical friction is much less efficient and therefore the detection probability is mostly determined by the lifetime of the QSO.

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