

A VLT/FORS2 NARROWBAND IMAGING SEARCH FOR MG II EMISSION AROUND $z \sim 0.7$ GALAXIES

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ABSTRACT

The mass and energy of galactic winds remain poorly constrained via QSO and “down-the-barrel” absorption line studies. One way to better constrain these parameters is to measure the spatial extent of the outflow in emission. We perform a VLT/FORS2 narrowband imaging of 5 star-forming galaxies at redshift $z = 0.67 - 0.69$ in the GOODS-S field as part of an effort to spatially resolve and constrain the radial extent of large-scale outflows traced by Mg II emission. These observations probe extended Mg II emission down to unmatched surface brightness limits of 5.74×10^{-19} ergs sec⁻¹ cm⁻² arcsec² (5σ). We do not detect any extended Mg II emission in any of our galaxies, thus placing strong upper limits on the strength of Mg II emission at projected distances $R_{\perp} = 13 - 24$ kpc from the sample galaxies. We do detect significant spatially-resolved Mg II absorption, revealing approximately constant absorption strengths across the galaxy disks. Our detection limits in concert with a previous studies Keck/LRIS spectra allow us to compare our observations to previous radiative transfer models of Mg II emission for isotropic and anisotropic dust/dust-free winds. Based on these comparisons, if our galaxies do host Mg II emission, then the predicted Mg II based on this modeling suggests that these winds are anisotropic and dusty for any scattered Mg II emission to lie below our 5σ detection limits.

Keywords: galaxies: evolution, galaxies: absorption lines

1. INTRODUCTION

Galactic winds play a critical role in regulating the star formation rates and stellar masses of galaxies (Werk et al. 2014); however, the physical mechanism that powers these galactic winds remains uncertain. Some possible mechanics have been proposed by theoretical studies that include thermal pressure from core collapse supernova, radiation pressure from starbursts and finally cosmic ray pressure (Larson 1974; Chevalier & Clegg 1985; Springel & Hernquist 2003; Sugahara et al. 2017). Additionally, the impact galactic winds have on their host galaxies (i.e, their mass and energy content) has remained difficult to constrain with observations.

An accurate picture of what types of galaxies host outflows comes from numerous absorption line studies of galaxies (Veilleux et al. 2005; Weiner et al. 2009; Martin et al. 2012; Rubin et al. 2014). Outflows are detected by measuring the blueshift (outflow) or redshift (inflow) of absorption transitions with respect to the host galaxy systematic velocity. Even though these probes can constrain: the radial velocity, column density and covering fraction of the flow, this method is weak in constraining the overall radial extent and provides no information on the morphology of the outflow.

One way of constraining the radial extent and morphology of the outflow is to trace winds in emission. This has been demonstrated using rest-frame optical transitions (i.e H α , [O III], etc) as tracers for winds around nearby starbursts (Matsubayashi et al. 2009; Veilleux et al. 2009; Tripp et al. 2011) as these transitions can trace the warm shock heated phase of the gas. A novel,

alternative transition potentially useful for tracing winds in emission is the Mg II $\lambda\lambda 2976, 2803$ doublet in the rest-frame ultraviolet (UV, [Weiner et al. 2009](#); [Kornei et al. 2013](#)). While most studies of winds using Mg II have focused on its absorption kinematics, [Rubin et al. \(2011\)](#) observed strong Mg II emission with a P-Cygni line profile in the spectrum of a strongly star-forming galaxy at redshift $z = 0.694$. This emission was the first detection of an outflow beyond the spatial extent of the galaxy continuum in the distant universe, with a minimum spatial extent > 7 kpc.

One proposed production mechanism for the galaxy’s observed Mg II line profile is photon scattering. In this mechanism, Mg II ions in the region of the wind closest to the observer will absorb continuum photons in the resonant transitions, λ_{2796} and λ_{2803} . Once these transitions are excited, they may only decay back to the ground state and if the optical depth of the gas is high, then the gas will resonantly trap the absorbed photons. Because the photons are absorbed in the rest frame of the gas, the absorption will be blueshifted relative to the galaxy’s systematic velocity. The Mg II ions in the section of the wind farthest from the observer will absorb and scatter photons that are redshifted relative to the front portion of the wind. Because the photons are redshifted, the photons will be able to travel freely towards the observer through the wind to produce emission at and redward of the systematic velocity of the galaxy ([Rubin et al. 2011](#), [Prochaska et al. 2011](#)). Since the first detection of Mg II emission by [Rubin et al. \(2011\)](#), another detection was reported by [Martin et al. \(2013\)](#), who observed Mg II emission that extends 12 – 18 kpc from a strongly star-forming galaxy at $z = 0.9392$. Mg II has also been studied in wide-field galaxy surveys conducted with Keck/DEIMOS and VLT/MUSE ([Erb et al. 2012](#), [Feltre et al. 2018](#)). These surveys, which include galaxies with redshifts $0.70 < z < 2.30$, find that Mg II may be detected in pure emission, pure absorption or with P-Cygni profiles. Detections of Mg II in emission were found to be associated with galaxies of lower stellar mass and bluer spectral slopes.

The diversity of these spectral profiles may be understood using radiative transfer modeling of galactic winds. [Prochaska et al. \(2011\)](#) have used this technique to predict spectra for the Mg II and Fe II* transitions for a variety of wind morphologies. The authors demonstrated that isotropic dust-free winds will conserve photon flux, thus predicting that blueshifted absorption lines should be accompanied by emission profiles of winds with similar equivalent widths (EW). Anisotropic winds, however, were demonstrated to exhibit significantly weaker emission at strengths proportional to the

angular extent (i.e., solid angle) of the wind. The scattered emission was additionally weakened by the inclusion of dust and the presence of a strongly-absorbing interstellar medium (ISM). Thus, spatially-resolved measurement of the surface brightness of this emission constrains not only the radial extent of the emitting material, but also its morphology and dust content.

In this paper, we present the first narrowband imaging of the Mg II transition around 5 star-forming galaxies located in the Great Observatories Origins Deep Survey south (GOODS-S) field at redshift $z \sim 0.7$. We use two filters: a line filter and a continuum filter. The line filter covers the Mg II doublet, while the continuum filter is offset from the line filter by ~ 47 Å. The resulting imaging in each filter has a total integration time of 10 hrs.

As opposed to 1D spectra, the narrowband imaging constrains the surface brightness, optical depth and radial extent of the wind. This narrowband imaging allows us to create the first ever high-S/N spatially resolved map of Mg II emission and absorption.

In Section 2 we describe our sample of GOOD-S galaxies, supplemental Keck/LRIS spectra, as well as our VLT/FORS2 observations, image reduction, and absolute flux calibration. We go into detail in our method of continuum subtraction in Section 3. The analysis of these data is presented in Section 4, including our methods for calculating surface brightness profiles and detection limits for each galaxy, as well as maps of Mg II equivalent widths. Section 5 presents results from our narrowband imaging analysis. We compare our SB detection limits to previous detections of extended Mg II emission, and compare our observations to predictions made by radiative transfer models in Section 6. We conclude this paper in Section 7. We adopt a Λ CDM cosmology with $h_{70} = H_0/(70 \text{ km s}^{-1} \text{ Mpc}^{-1})$, $\Omega_M = 0.3$, and $\Omega_\Lambda = 0.7$. In this cosmology, $1''$ is ≈ 7 kpc at $z \sim 0.7$.

2. OBSERVATIONS AND DATA REDUCTION

2.1. Sample Selection

Our target galaxies were selected from a Keck/LRIS survey of UV absorption lines in ≈ 100 objects having redshifts $0.3 < z < 1.4$ and $B_{AB} < 23$ in fields with deep *HST*/ACS imaging ([Rubin et al. 2014](#)). In particular, this parent survey targeted galaxies in a total of nine Keck/LRIS pointings located in both of the GOODS fields ([Giavalisco et al. 2004](#)) and the AEGIS survey field (the Extended Groth Strip; [Davis et al. 2007](#)). In inspecting the redshift distribution of the portion of this sample observable from the Southern Hemisphere, we uncovered a narrow peak of nine galaxies in

the interval $0.66 < z < 0.68$. This peak is in fact the global maximum of the distribution, as all other bins of width $\Delta z = 0.02$ have at most four galaxies. Moreover, there are two narrow interference filters available on VLT/FORS2 centered at $\lambda \sim 4675$ and 4722 \AA which cover the Mg II $\lambda\lambda 2796, 2803$ transition in precisely this redshift interval. We selected our final sample of five of these galaxies at $0.66 < z < 0.68$ to be close on the sky such that they could be imaged in a single $7' \times 7'$ FORS2 pointing. We show color *HST*/ACS images of these objects in Figure 1.

The absorption line modeling presented in Rubin et al. (2014) indicates that these five galaxies are driving strong outflows traced by Mg II with velocities $\sim 150 - 420 \text{ km s}^{-1}$ and equivalent widths (EW) $\sim 2 - 3 \text{ \AA}$. Modeling of the galaxy broad-band spectral energy distributions obtained from multi-wavelength ancillary imaging data yields star formation rates (SFR) ranging from ~ 4 to $40 M_{\odot} \text{ yr}^{-1}$ and stellar masses in the range $\log M_*/M_{\odot} \sim 9.9 - 11.0$. The properties and nomenclature of the sample, as well as precise target coordinates taken from Rubin et al. (2014) are listed in Table 1.

2.2. VLT/FORS2 Observations

Our narrow-band imaging data were taken in service-mode using the FORS2 instrument on the VLT 8.2m telescope Antu between October 2012 and February 2013. We used two narrowband filters, HeII+47 and HeII/3000+48, that have peak transmission at wavelengths that correspond to the Mg II doublet lines at our sample redshift of $z \sim 0.7$ (see Table 2). The filter transmission curves are plotted along with each galaxy’s spectrum in Figure 2. In the following, we will often refer to the HeII+47 filter as the “line” or Mg II filter and the HeII/3000+48 filter as the “continuum” filter.

FORS2 has a native pixel scale of $0.125'' \text{ pixel}^{-1}$ and a field of view of $7' \times 7'$. The data were taken with the CCD binned 2×2 , yielding a pixel scale of $0.25'' \text{ pixel}^{-1}$. Images of three pointings offset by $0.25'$ East/West were obtained, with individual exposure times of $\approx 1000 \text{ sec}$. A total of 38 exposures were taken in each filter. Our observations were carried out under photometric and thin cloud conditions (program ID: 090.A-0427A). The seeing values, given in the header of each image, are derived from zenith observations at 0.5 microns with the Paranal differential image motion monitor (DIMM, Sarazin & Roddier 1990). The initial DIMM values are then corrected to the airmass and wavelengths of the science observation. Finally, a first order correction is then applied for appropriate for a 8m telescope is applied (see Figure 3). The median seeing for the images is $\sim 0.8''$

Summing the individual exposure times for each filter results in a combined exposure time of 10.0 hours each for the HeII+47 and HeII/3000+48 filters.

2.3. Supplemental Keck/LRIS spectra

In addition to VLT imaging, in the present analysis we utilize galaxy spectra taken from the Rubin et al. (2014) Keck I Low Resolution Imaging Spectrometer (LRIS) program. A $0.9''$ slit width was used for all slitmasks and the median FWHM resolution for the spectra is 274 km s^{-1} at $\lambda_{\text{rest}} \approx 2800 \text{ \AA}$ and 286 km s^{-1} at $\lambda_{\text{rest}} \approx 2600 \text{ \AA}$ (see Figure 2). The spectral coverage of these data extends from ~ 3200 to 8000 \AA .

2.4. Image Reduction

The data were reduced using custom routines written in *Python*. The images were first corrected by subtracting and removing the overscan region of the CCD. Then the images were bias-subtracted and flat-fielded using twilight flats. To improve the flat-fielding essential for detecting faint extended emission across the fields, we further correct for the illumination patterns using night-sky flats. The night-sky flats were produced by combining the unregistered science frames with an average sigma-clipping algorithm after masking out all the objects, and any bad pixels. Each individual image was cleaned of cosmic rays and bad pixels by utilizing the *L.A. Cosmic* algorithm (van Dokkum 2001). The astrometry solutions were calculated via Astrometry.net (Lang et al. 2010), and yield a standard deviation in the galaxy coordinates of $\sigma \approx 0.10''$. Before image stacking, we ran each frame through *SExtractor* (Bertin et al. 1996) in order to create a root mean square (RMS) map of each science image.

The final stacked image for each filter is obtained using *SWarp* (Bertin et al. 1996). Each individual frame is first sky-subtracted using a background mesh size of 256 pixels which is approximately $64''$. We chose the mesh size to be large enough such that the extended emission is not mistakenly subtracted (Arrigoni Battaia et al. 2015). The frames, after background-subtraction, are resampled onto a common astrometric solution using a *Lancosz3* interpolation kernel. The images are weighted by the night-flat image and then average-combined in order to increase the signal-to-noise of any Mg II emission. Additionally, *SWarp* generates stacked RMS images by propagating the error images for each science frame. Our final stacked images in each filter are shown in Figure 4 with the target galaxies indicated.

2.5. Absolute Flux Calibration

We acquired observations of the standard star GD50 from archival ESO calibration imaging at 4 indepen-

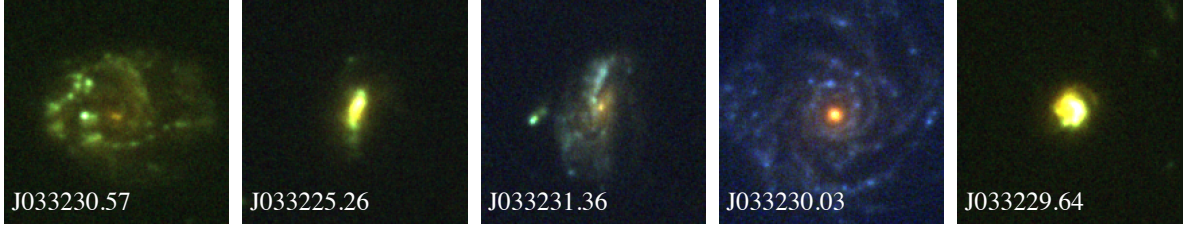


Figure 1. Color imaging of our sample galaxies in the *HST*/ACS F435W, F606W, and F775W filters obtained as part of the GOODS survey (Giavalisco et al. 2004). Each image is $5'' \times 5''$ (or about $40 \text{ kpc} \times 40 \text{ kpc}$).

Table 1. Properties of the 5 galaxies in our sample as estimated in Rubin et al. (2014). The EW (in the observed frame) include both components of the Mg II doublet and were determined from analysis of the supplemental Keck/LRIS spectra.

Object ^a	R.A.	Dec	z	$\text{SFR}(M_{\odot} \text{ yr}^{-1})$	$\log M_{*}/M_{\odot}$	$\text{EW}_{\text{obs}}(\text{\AA})$	τ_V^b
J033225.26-274524.0 (J.26)	03:32:25.26	-27:45:23.9	0.6660	$9.1^{+1.3}_{-3.7}$	$9.86^{+0.05}_{-0.04}$	7.539 ± 0.354	$1.227^{+1.54}_{-0.20}$
J033231.36-274725.0 (J.36)	03:32:31.35	-27:47:24.9	0.6669	$10.5^{+1.7}_{-1.6}$	$10.02^{+0.03}_{-0.03}$	5.835 ± 0.493	$1.377^{+0.60}_{-0.23}$
J033230.03-274347.3 (J.03)	03:32:30.03	-27:43:47.2	0.6679	3.8^{+0}_{-0}	$10.98^{+0.01}_{-0.0}$	12.794 ± 1.710	$0.297^{+0.0}_{-0.0}$
J033229.64-274242.6 (J.64)	03:32:29.64	-27:42:42.5	0.6671	$40.5^{+8.2}_{-12.1}$	$10.30^{+0.07}_{-0.03}$	13.239 ± 0.263	$3.897^{+1.15}_{-0.93}$
J033230.57-274518.2 (J.57)	03:32:30.56	-27:45:18.2	0.6807	$12.6^{+1.7}_{-2.1}$	$10.48^{+0.03}_{-0.07}$	6.106 ± 0.370	$1.262^{+1.23}_{-0.40}$

^a Name in parenthesis to be used as shorthand for the objects throughout the remainder of the paper.

^b Total V-band optical depth of dust.

Table 2. Filter properties and exposure times of the VLT/FORS2 observations. The width of the transmission curves ($\Delta\lambda$) are calculated by convolving the transmission curve with the total wavelength range of the filter. These values differ slightly from those reported by the European Space Observatory.

Filter	$\lambda_{\text{eff}}(\text{\AA})^a$	$\Delta\lambda(\text{\AA})$	N^b	$T(s)^c$	S^d
HeII+47	4675.21	50.11	38	35,959	2.45
HeII/3000+48	4722.46	44.82	38	36,937	2.40

^a λ_{eff} is the effective wavelength of the filter transmission curve.

^b Total number of images.

^c Total exposure time.

^d S , the sensitivity of the filter, is in units of $10^{-17} \text{ ergs counts}^{-1} \text{ cm}^{-2}$.

dent epochs. Performing aperture photometry at each epoch and airmass, we calculated the atmospheric extinction coefficients, k , to be 0.181 magnitudes for the HeII/3000+48 filter and 0.190 magnitudes for the HeII+47 filter. We perform absolute flux calibration using the methods of Jacoby et al. (1987). We first convolve the spectral energy distribution of the standard star, $F(\lambda)$ in $\text{ergs sec}^{-1} \text{\AA}^{-1} \text{ cm}^{-2}$, with that of the known transmission curve of the filter, $T_i(\lambda)$. This yields F_i , the total observable flux in each bandpass filter i with units of $\text{ergs sec}^{-1} \text{ cm}^{-2}$:

$$F_i = \int F(\lambda) T_i(\lambda) d\lambda.$$

It is not uncommon to assume that $F(\lambda)$ is constant over the small width of the filter. However, we calculate the integral numerically. The system sensitivity, including the defects of the telescope optics and detector response is then given by

$$S_i = \frac{F_i}{C 10^{k_i A}},$$

where k_i is the extinction in magnitudes per airmass, A is the airmass for each individual exposure, C is the measured count rate of the standard star and S_i is in units of $\text{ergs counts}^{-1} \text{ cm}^{-2}$. Before image co-addition, each science image is corrected for atmospheric extinction by multiplying each frame by $10^{k_i A}$. Next, the image is divided by the exposure time, effectively putting the image in units of counts per sec. After co-addition, the images are then multiplied by the appropriate sensitivity factor S_i . This puts the final images in the appropriate flux units, $\text{ergs sec}^{-1} \text{ cm}^{-2}$.

3. IMAGE SUBTRACTION

We have two goals for our study: (1) assess the surface brightness of line emission in the Mg II transition in and around each target galaxy; and (2) spatially resolve the morphology of the strong Mg II absorption observed against the galaxy continua. To reach both of these goals, we must perform accurate subtraction of the continuum flux of each object from the filter covering the targeted line emission. For four of the five galaxies in

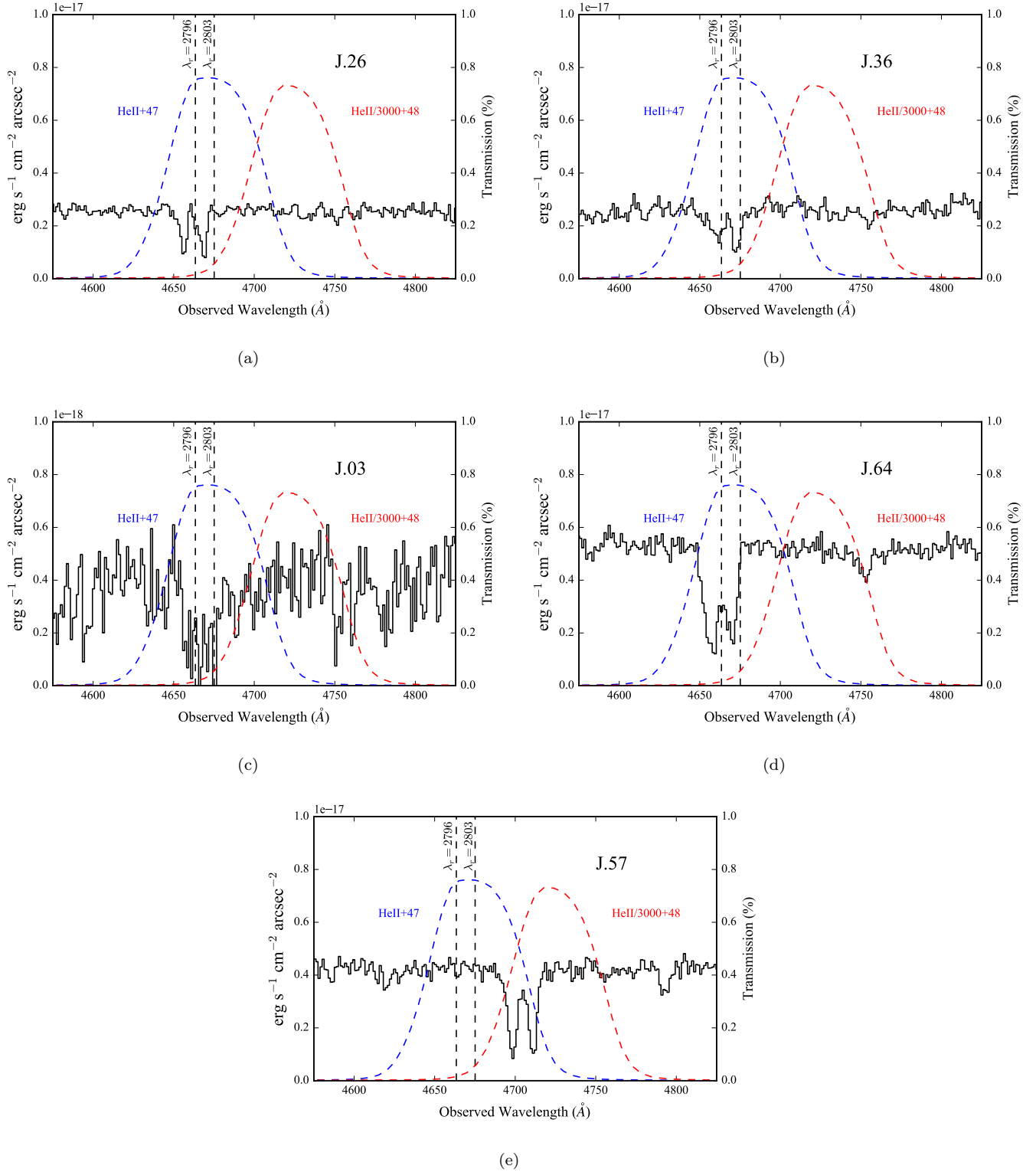


Figure 2. Keck/LRIS spectra of the sample galaxies and the transmission curves of the filters HeII+47 (blue dashed line) and HeII/3000+48 (red dashed line). The left-hand axis is in units of flux density and the right-hand axis is the percentage of light transmitted by the filter at each wavelength. Vertical dashed lines indicate the wavelengths of the redshifted Mg II doublet. The Mg II doublet falls fortuitously at the central wavelength of the HeII+47 filter for the galaxies shown in panels (a) through (d).

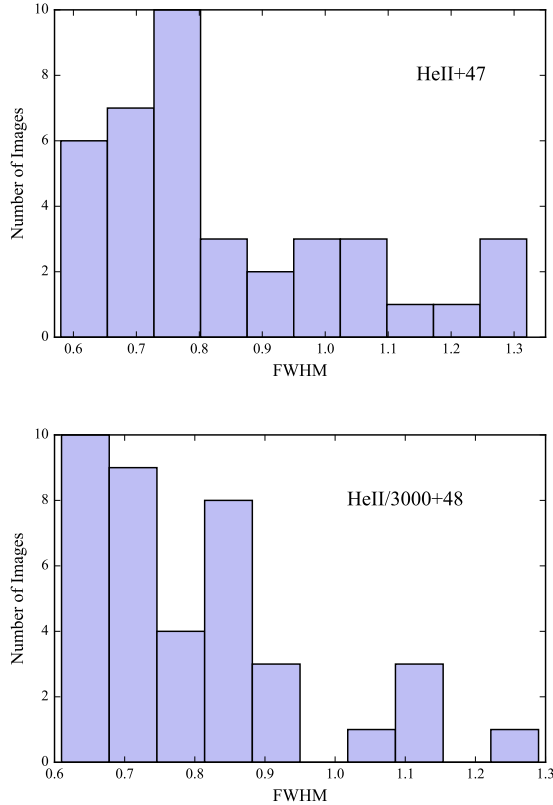


Figure 3. Top: Distribution of seeings for the 38 HeII+47 images. **Bottom:** Same HeII/3000+48 images. The median seeing conditions for the images in both filters is $\sim 0.8''$. The seeing conditions were calculated from DIMM measurements which are provided by ESO in the header of each science image.

our sample, the HeII+47 image includes both line and continuum emission, and the HeII3000+48 image provides a high S/N measurement of the continuum only $\approx 30 \text{ \AA}$ redward of the line emission in the rest-frame. The galaxy excluded from our analysis is instead used as a check on the quality of our continuum subtraction for reasons detailed in Section 5.1.1.

3.1. Spectral Correction

In preparation for continuum subtraction, we first consider whether the continuum level of each galaxy spectrum changes significantly over the passbands of our two filters. We use the supplementary spectra from Rubin et al. (2014) to fit the continuum and determine the spectral slope of each galaxy. We use the interactive fitting routine *lt_continuumfit* from the *linetools* package (Prochaska 2016)¹ to fit the continuum. We then find

the total continuum flux in each filter by convolving the fitted continuum with each filter’s transmission curve. Next, we take the ratio of both integrated totals, as the ratio will indicate the scaling factor needed to correct our flux measurements prior to continuum subtraction. Comparing these ratios between each galaxy, we find that they are effectively the same, within 0.1%, with a value of 1.118. This value is equal to the ratio between the FWHM of the filter transmission curves, allowing us to conclude that the slope of the spectrum of each galaxy is flat, and hence that the continuum level measured in the off-line filter provides an accurate measure of the continuum contribution to the on-line filter flux.

3.2. Continuum Subtraction

To properly continuum-subtract the image taken with the Mg II filter, we follow a prescription given by Arigoni Battaia et al. (2015). We first determine the continuum flux density from the continuum filter,

$$f_{\text{cont}} = \frac{F_{\text{cont}}}{\Delta\lambda_{\text{cont}}}, \quad (1)$$

where F_{cont} and $\Delta\lambda_{\text{cont}}$ are the observed flux per pixel of the continuum image and the transmission FWHM of the continuum filter, respectively. With f_{cont} it is then possible to calculate the flux of any excess emission, F_{line} :

$$F_{\text{line}} = F_{\text{MgII}} - f_{\text{cont}}\Delta\lambda_{\text{MgII}} \quad (2)$$

where F_{MgII} and $\Delta\lambda_{\text{MgII}}$ are the observed flux per pixel in the Mg II filter and the transmission FWHM of the Mg II filter. The continuum subtracted images of each galaxy are shown in Figure 5. The continuum subtracted image has uniform background and no obvious signatures of emission.

4. ANALYSIS

4.1. Surface Brightness Profiles and Limits

In order to test for the presence of Mg II emission, we perform aperture photometry on the continuum subtracted images using the python library *Photutils*. We choose annuli with a radial thickness of 1 pixel or $0.25''$, such that, $r_{\text{inner}} = r_{\text{outer}} - 1$ (in pixels). Each annulus is centered on the flux-weighted centroid of the galaxy. By dividing the summed flux in each annulus by the area in arcseconds we produce surface brightness (SB) profiles for each galaxy. These profiles are shown in Figure 6.

The error in the SB is determined from the stacked RMS images of each object. We adopt annuli that are identical to the annuli used to find the SB profiles for each galaxy. To calculate the variance inside each annulus, we sum the RMS pixel values in quadrature, then divide by the area of each annulus.

¹ <https://github.com/linetools/linetools>

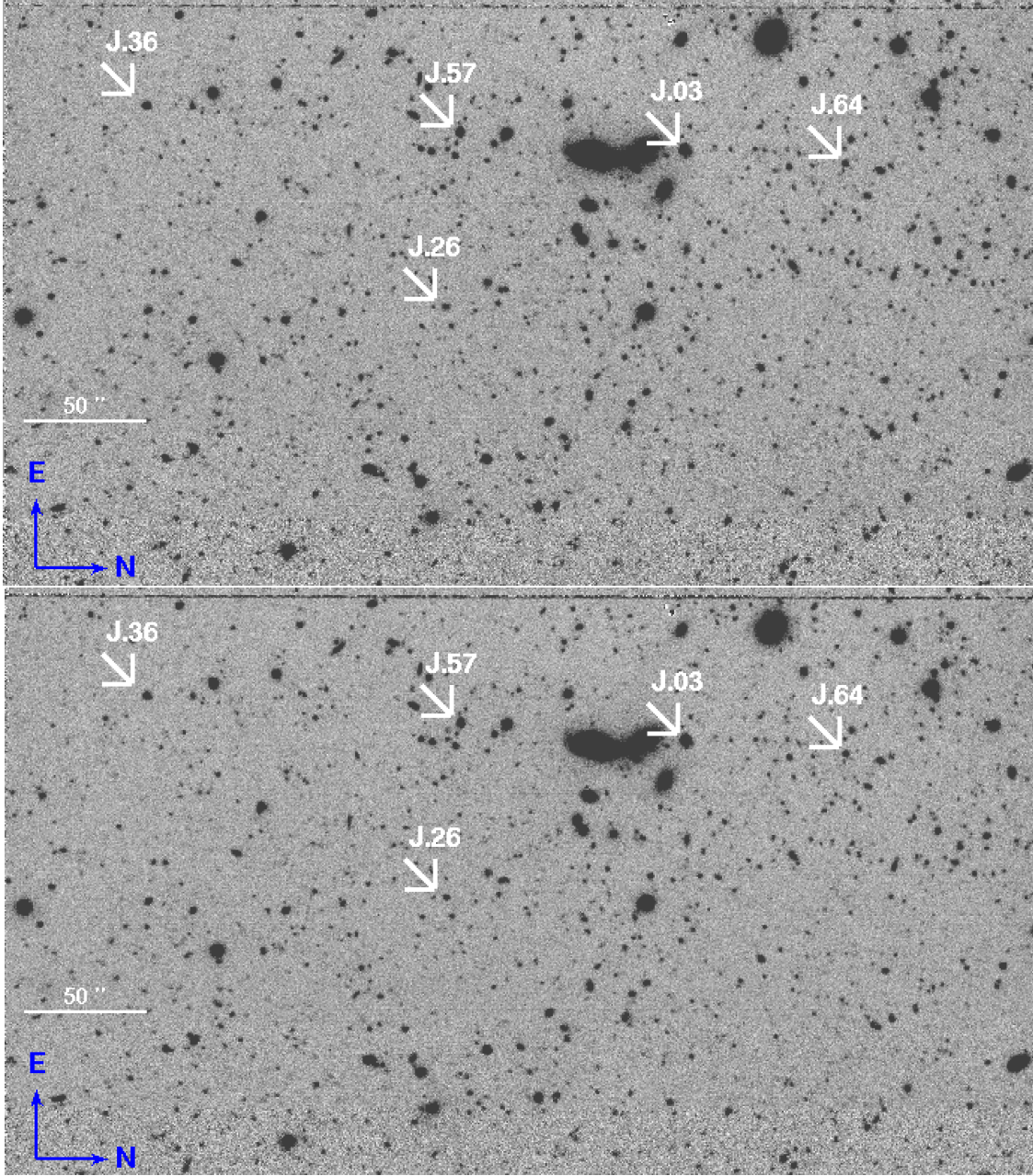


Figure 4. Top: Stacked HeII+47 image of the galaxy sample. Bottom: Stacked HeII3000+48 image of the same pointing. The exposure time of each image is ≈ 10 hours. Each image shows half of the total FOV of $7' \times 7'$ which contains the full sample of galaxies (indicated by the white arrows). East is up and North is right.

To calculate the 1σ SB limit we follow the procedure of Arrigoni Battaia et al. (2015). We first mask out all the sources, their associated extended halos, and edge noise in both the HeII+47 and HeII/3000+48 images. We then calculate the RMS of the background in randomly-placed $1''$ apertures. We convert these RMS values to SB limits per 1 arcsec^2 aperture. We find that the 1σ detection limits (SB_1) are $6.332 \times 10^{-19} \text{ ergs sec}^{-1} \text{ cm}^{-2} \text{ arcsec}^2$ and $5.808 \times 10^{-19} \text{ ergs sec}^{-1} \text{ cm}^{-2} \text{ arcsec}^2$ in the HeII/3000+48 and HeII+47 filters, respectively.

With the 1σ detection limit, SB_1 , determined for the continuum+Mg II (HeII+47) image, we define a thicker (or “extended”) annulus to be used to search for any extended Mg II emission. This extended annulus has an inner radius approximately the size of the SB_1 isophotal contour for each galaxy. The outer radius is chosen to be the inner radius plus 5 pixels. Measurements of Mg II SB will probe projected distances of 13, 18, 21 and 24 kpc from the centers of each target. The resulting SB measurements are shown in Figure 6.

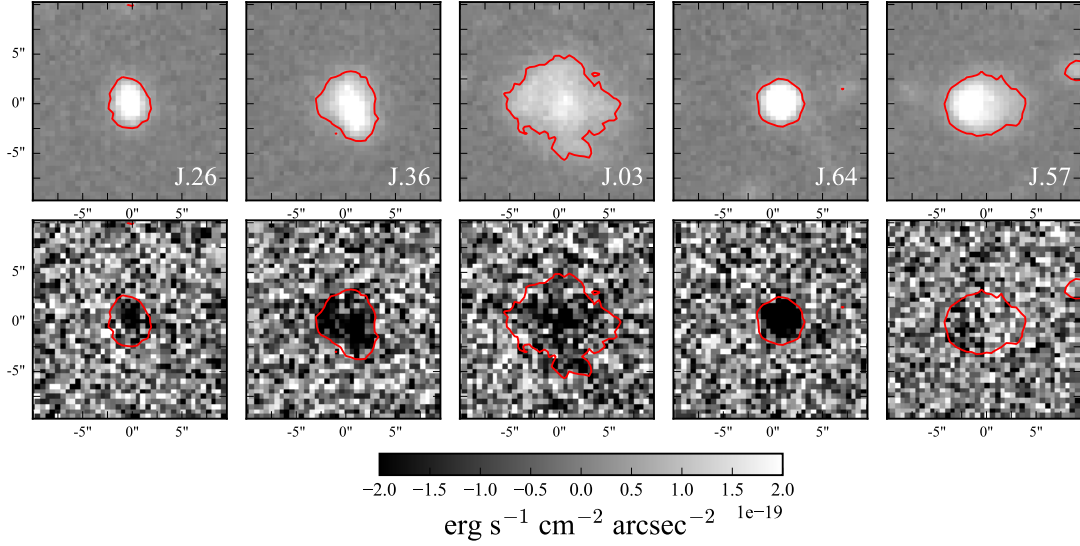


Figure 5. $10'' \times 10''$ (or about $70 \text{ kpc} \times 70 \text{ kpc}$) images of each galaxy in our sample. Top row: Continuum surface brightness in $\text{ergs s}^{-1} \text{ cm}^{-2} \text{ arcsec}^{-2}$ measured in the HeII/3000+48 filter. Bottom row: Continuum-subtracted Mg II surface brightness. Absorption can be seen in 4 of 5 galaxies. The red contours represent the outline of the 1σ surface brightness limit in the HeII+47 image, defined in Sec. 4.1. The colorbar shows the scaling used for the Mg II images.

The sensitivity required to detect an extended source depends on its size, as one can reach lower SB levels by spatially averaging over large apertures. In the ideal case of perfect sky subtraction and continuum subtraction, the 1σ SB limit for an extended source is $SB_1/\sqrt{A_{\text{src}}}$, where A_{src} is the area in arcsec^2 and SB_1 is the surface brightness limit per 1 arcsec^2 aperture. In practice, the actual detection limits are affected by systematics from imperfect subtraction. Therefore, we empirically determine the limits as follows. We mask all the artifacts and sources in the continuum-subtracted images. Next, we generate apertures with sizes similar to our extended annulus $\sim 20 \text{ sq.arcsec}$, place them at random, and extract the fluxes, F_{src} , within these apertures.

In the ideal case that the sky and continuum are perfectly subtracted the value of $F_{\text{src}}/\sigma_{\text{src}}$, where $\sigma_{\text{src}} \equiv SB_1/\sqrt{A_{\text{src}}}$, from many random apertures should follow a Gaussian distribution with unit variance. The distribution of $F_{\text{src}}/\sigma_{\text{src}}$ for these apertures is shown in Figure 7. We calculate the variance and mean of the distribution and find that the variance of the distribution is $\sigma'_{\text{src}} = 1.1$ and implies that the SB detection limit for our continuum-subtracted image is broader than σ_{src} by 10%. We thus adopt $F_{\text{limit}} \equiv \sigma'_{\text{src}}$ as the 1σ upper limit on the total line flux of extended Mg II emission. The SB_{limit} is then $F_{\text{limit}}/A_{\text{src}}$. The values of F_{src} and the $5SB_{\text{limit}}$ for each galaxy are listed in Table 3.

Table 3. Significance of extracted flux and detection limits

Object	$F_{\text{src}}(\text{Mg II})^a$	$5SB_{\text{limit}}^b$	Area ^c
J033225.26-274524.0	2.44(0.92)	6.51	21
J033232.36-274725.0	-1.40(-0.53)	6.51	21
J033230.03-274347.3	-5.23(-1.75)	5.74	27
J033229.64-274242.5	1.23 (0.44)	6.22	26
J033230.57-274518.2	-2.53 (-1.00)	6.81	18

^a Mg II flux is in $10^{-18} \text{ ergs sec}^{-1} \text{ cm}^{-2}$. The value in the parenthesis is the statistical significance with respect to σ_{src} .

^b Limits are in $10^{-19} \text{ ergs sec}^{-1} \text{ cm}^{-2} \text{ arcsec}^{-2}$.

^c Area of the extended annulus in sq.arcsec .

In order to show that our detection limits are reasonable, we simulate emission with varying intensities. The results of this exercise are shown in Figure 8. For each galaxy, we assign our simulated emission a constant surface brightness corresponding to 1, 3, 5, 10 and 20 times the 1σ SB_{limit} inside the largest annulus used (i.e., the extended annulus). Additionally, we assume Gaussian noise with 1σ equal to $1SB_{\text{limit}}$. After placing the simulated emission around the galaxy, we subtract the continuum in the same manner as explained in Section 3.2.

We then construct a smoothed χ image following the techniques in Hennawi & Prochaska (2013) and Arrigoni Battaia et al. (2015). First, we smooth the continuum-subtracted image:

$$I_{\text{smooth}} = \text{CONVOLVE}[\text{NB-CONTINUUM}], \quad (3)$$

4.2. Test of Surface Brightness Limits

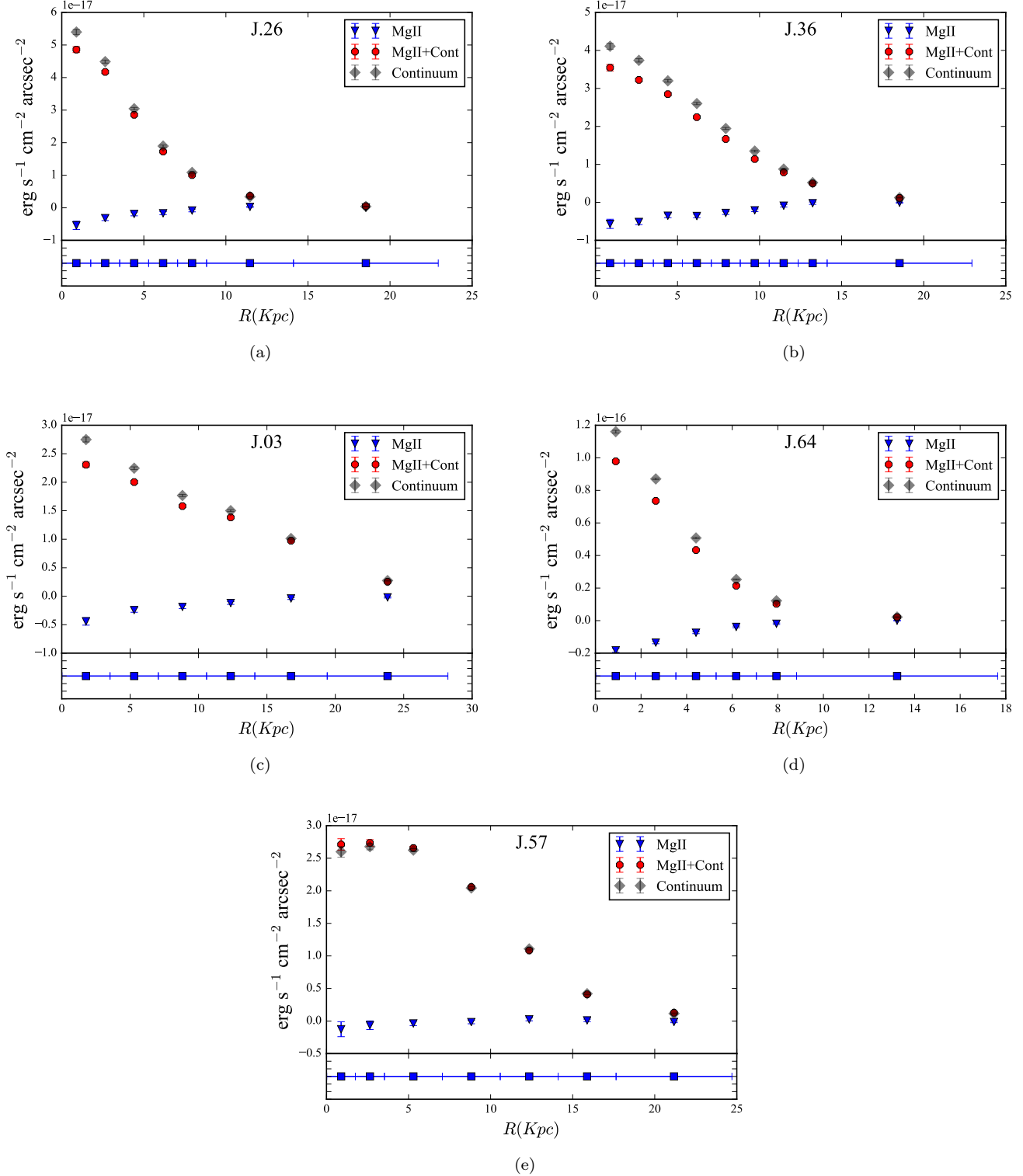


Figure 6. SB profiles for our sample galaxies. (Top panel) Continuum SB profile (black) measured for each galaxy. The green points show the Mg II + continuum SB measured for the galaxy in the pre-continuum subtracted image. The blue points show the Mg II line SB measured for the galaxy in the continuum subtracted line emission image. The profile exhibits SB decrements from Mg II absorption. Photometry was performed in circular annuli. (Bottom panel) The vertical hashes in the bottom show the inner and outer radius of each annulus in kpc. Distance from the center of the galaxy (x-axis) is computed using the average value of the inner and outer radii of each annulus.

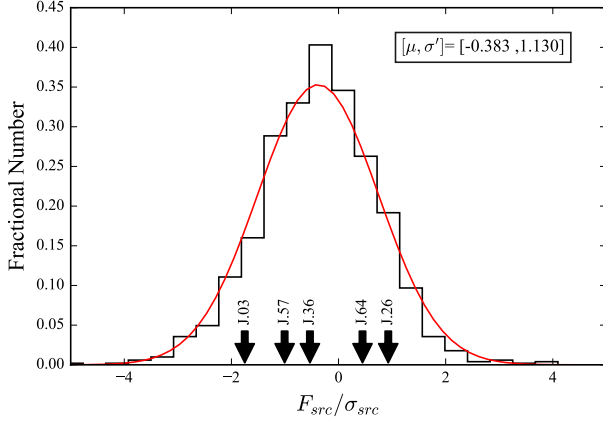


Figure 7. Normalized distribution of $F_{\text{src}}/\sigma_{\text{src}}$ values for random circular annuli placed on the continuum-subtracted line image. F_{src} is the total flux within an aperture and σ_{src} is the expected 1σ flux limit in the ideal case of perfect sky and continuum subtraction, i.e. $SB_1\sqrt{A_{\text{src}}}$. The black arrows point to the statistical significance of the flux inside the “extended annulus” of each galaxy.

where the CONVOLVE operation indicates convolution of the images with a Gaussian kernel with FWHM=1.5 pixels. Next, we computed the sigma image (σ_{smooth}) for the smoothed image (I_{smooth}) by propagating the noise image of the unsmoothed data:

$$\sigma_{\text{smooth}} = \sqrt{\text{CONVOLVE}^2[\sigma_{\text{unsmooth}}^2]}, \quad (4)$$

where the CONVOLVE² operation indicates the convolution of the variance image with the square of the Gaussian kernel. The smoothed χ image is defined by

$$\chi_{\text{smooth}} = \frac{I_{\text{smooth}}}{\sigma_{\text{smooth}}}. \quad (5)$$

This χ_{smooth} image aids in recognizing the presence of extended Mg II emission.

Figure 8 shows the χ_{smooth} images for the 5 levels of simulated Mg II emission. We also include the χ_{smooth} image of each galaxy without any simulated emission (in the left most column). The galaxies are outlined by a black isophotal contour corresponding to $1SB_1$ and the simulated emission is contained inside the extended annulus surrounding each contour. The χ_{smooth} images confirm that we should be able to detect extended Mg II emission down to a conservative level of $5SB_{\text{limit}}$. Note again that the SB_{limit} does indeed take into account the systematics from imperfect continuum subtraction.

4.3. Equivalent Widths

Here we derive an expression to calculate the equivalent width (EW_{MgII}) of any absorption or emission features observed in our narrow-band imaging. Starting

from the expression for EW used in the context of spectroscopy,

$$EW_{\lambda} = \int \left(1 - \frac{f_{\lambda}}{f_{\text{cont}}}\right) d\lambda \quad (6)$$

we begin by dividing Eq 2 by the flux density of the continuum and the FWHM of the on-line filter,

$$\frac{F_{\lambda}}{f_{\text{cont}}\Delta\lambda_{\text{MgII}}} = \frac{F_{\text{MgII}}}{f_{\text{cont}}\Delta\lambda_{\text{MgII}}} - 1. \quad (7)$$

Next, we rearrange the above expression such that we produce the argument of the integrand in Eq. 6 on the right hand side,

$$-\frac{F_{\lambda}}{f_{\text{cont}}\Delta\lambda_{\text{MgII}}} = 1 - \frac{f_{\text{MgII}}}{f_{\text{cont}}}. \quad (8)$$

We then approximate the integration in Eq. 6 by multiplying the integrand above by the FWHM of the on-line filter $d\lambda = \Delta\lambda_{\text{MgII}}$,

$$-\frac{F_{\lambda}}{f_{\text{cont}}} = \left(1 - \frac{f_{\text{MgII}}}{f_{\text{cont}}}\right)\Delta\lambda_{\text{MgII}}; \quad (9)$$

such that

$$EW_{\text{MgII}} = -\frac{F_{\text{MgII}}}{f_{\text{cont}}}. \quad (10)$$

Using the above equation along with the continuum and continuum-subtracted images, we produce images of the observed-frame EW_{MgII} . The observed-frame EW_{MgII} images are displayed for each galaxy in Figure 11 and show only the EWs within the 1σ SB_1 contours of the corresponding Mg II images (prior to continuum subtraction).

To compare our map of EW_{MgII} to the values measured from the Keck/LRIS spectra, we place 0.9 arcsec wide apertures on top of each galaxy. The width and angle of the apertures replicate the orientation of the slits used to obtain the spectra. Next, we determine which pixels lie outside the 1σ SB_1 contours and set their values to zero. Outside this contour, the EW_{MgII} values become poorly constrained due to the lack of S/N in the continuum. We then apply a $S/N = 1.5$ cut to the continuum image and create a histogram to show the distribution of EW_{MgII} inside the contour. The histograms are shown in Figure 11. Having removed pixels with low-S/N continuum values from our images, we compute the mean equivalent width inside the Keck/LRIS apertures. These values are summarized in Table 4.

To assess the morphology of the Mg II absorption, we determine the distance of each pixel from the center of each galaxy in kiloparsecs. We plot the EW_{MgII} over this projected distance for each galaxy in Figure 11. Although some of the plots suggest a slight upward trend in the values of the EW_{MgII} with increasing radii, we

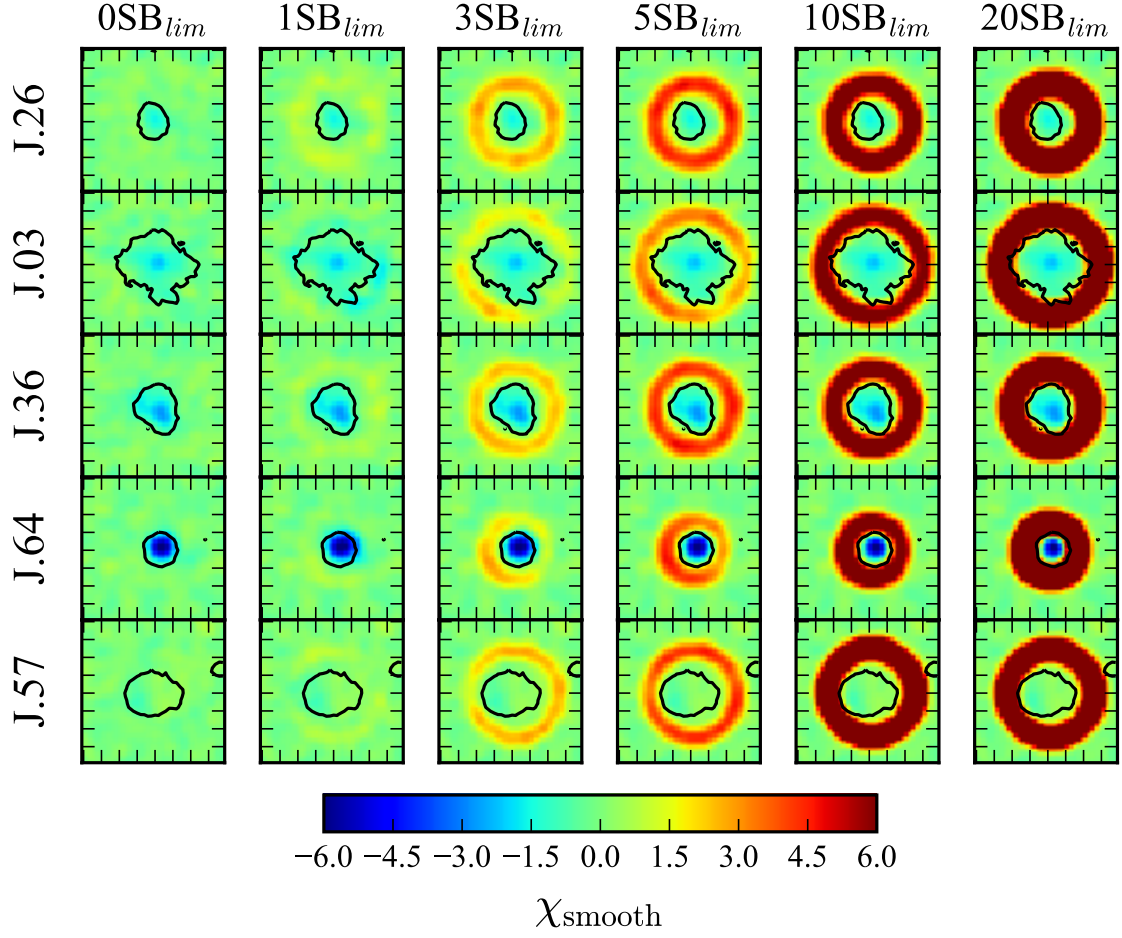


Figure 8. Postage stamp χ_{smooth} images for the Mg II emission of the 5 galaxies in our sample. Every galaxy is placed in the same row in each column. The columns show simulated emission, with levels of 0, 1, 3, 5, 10, and 20 times SB_{lim} . Each postage stamp has a size of $5'' \times 5''$ (corresponding to $35 \text{ kpc} \times 35 \text{ kpc}$ at $z \sim 0.70$). Each stamp shows the galaxy along with the same isophotal contour used in previous figures.

cannot be confident in this trend because of the large scatter. To better visualize the data and test the significance of the trend, we bin the data radially in 3-5 kpc-wide bins. For example, in Figure 11(c), the EW_{MgII} from absorption of J.03 extends out to ~ 25 kpc. The EW_{MgII} values are binned in 5 kpc increments. We calculate the mean and scatter of the EW_{MgII} in each bin and show these values in Figure 12.

5. RESULTS

5.1. Spatially Resolved Maps of Mg II Absorption

In this section we discuss the details of the absorption detected in our SB profiles as well as compare our EW_{MgII} measurements to those measured in the supplemental Keck/LRIS spectra.

5.1.1. Mg II absorption in surface brightness profiles

Although we do not detect any extended Mg II emission, we do observe a decrement of flux, SB_{abs} , in the SB profiles of 4 out of 5 galaxies in our sample. Figure 6 shows the SB profiles for the 5 galaxies in our sample. Absorption from Mg II ions is prevalent in the profiles between 10 - 25 kpc, decreasing radially outward from the maximum absorption at the center of the galaxies. In Table 4 we report for the galaxies J.26, J.36 and J.03 a maximum decrement in the SB profile due to absorption $\sim -5.00 \times 10^{-18} \text{ erg s}^{-1} \text{ cm}^{-2}$ per sq.arcsec. Additionally, we report for galaxy J.64 a SB_{abs} one order of magnitude greater than the previous galaxies with a value of $-18.2 \pm 0.127 \times 10^{-18} \text{ erg s}^{-1} \text{ cm}^{-2}$ per sq.arcsec. Finally, for galaxy J.57, SB_{abs} is strongest at the center of the galaxy with a SB value of $-1.25 \pm 1.15 \times 10^{-18} \text{ erg s}^{-1} \text{ cm}^{-2}$ per sq.arcsec. However, the measured SB of this decrement including error is consistent with measuring zero absorption. The measurement of zero absorption for J.57 was expected. Figure 2 shows the Keck/LRIS spectrum of J.57 and the transmission curves of the filters HeII+47 and HeII/3000+48. The Mg II transitions of the galaxy are equally sampled by both of our filters. When we subtract the continuum image from the Mg II image we are effectively subtracting the Mg II as well. This measurement suggests that the quality of our subtraction is satisfactory.

The surface brightness profiles presented in Figures 6 do not exhibit any signs of extended Mg II emission.

5.1.2. Morphology of MgII Absorption

Figures 11 show the images, distributions and radial projections of Mg II EWs. We have zeroed out any values that lie outside the SB_1 contours for each galaxy. We also impose a signal-to-noise cut, only including EWs where the continuum S/N is greater than 1.5. The mean

EW is computed for all pixels inside each Keck/LRIS aperture, defined in Sec. 4.3, and the resulting values are summarized in Table 4. Comparing our narrowband EWs with those measured from the spectra, we find agreement to within a $1.6\text{--}4.6\sigma$ for galaxies J.26 and J.36, and more statistically significant differences for galaxies J.03 and J.64. We discuss possible causes for these differences below.

Figure 2(c) shows the Keck/LRIS spectrum of galaxy J.03. The continuum observed near the Mg II transition is noisy, compared to the spectra of the rest of the sample. Since the value of the EW is dependent on the level of the continuum present near the Mg II transition, it may well be that our choice of continuum level in calculating the EW from the spectrum might be higher relative to the continuum level implied by our narrow-band image.

Figure 2(d) shows the Keck/LRIS spectrum of galaxy J.64. This galaxy is the brightest galaxy in the sample, and also exhibits the highest-velocity wind. This shifts the Mg II absorption profile toward the blue end of the HeII+47 transmission curve, which could cause the flux in this filter to be weighted toward the continuum level and the absorption signal to be underestimated.

As demonstrated in our radially-binned EW profiles, a majority of the galaxies exhibit large scatter in the binned EWs out to larger radii. To better understand the significance of these trends, we construct a plot that compiles the mean EWs for all the galaxies, shown in Figure 12. To take into account the varying size of the galaxies, we normalize the distances by the approximate radius of the SB_1 contour for each galaxy.

Upon inspection of Figure 12, we see that the galaxies exhibit no statistically significant trend in the mean absorption EW as a function of radius inside our 1SB_1 isophotal contour, which suggests that the covering fraction is approximately constant across the surface.

5.2. Limits on Mg II Emission

For our sample of galaxies, we are sensitive to extended emission at minimum projected distances of 13, 18, 21 and 24 kpc, respectively, from the centers of each target.

We do not detect any significant Mg II emission at these distances around any of our target galaxies. The χ_{smooth} images shown in Figure 8 confirm this. A comparison of the simulated emission with the χ_{smooth} image of the original continuum-subtracted image, shown in the first column, similarly suggests that we do not detect any extended Mg II emission. We thus place 5σ upper limits on Mg II emission for each galaxy in the sample, summarized in Table 3. The most sensitive de-

Table 4. Properties of Mg II Absorption

Object	SB_{abs}^a	R^b	$EW_{obs}(\text{\AA})^c$	$EW_{obs}(\text{\AA})^d$
J.26	-5.39 ± 1.28	8	3.51 ± 0.80	7.53 ± 0.35
J.36	-5.64 ± 0.121	15	7.69 ± 1.02	5.83 ± 0.49
J.03	-4.41 ± 0.622	21	5.39 ± 0.71	12.7 ± 1.71
J.64	-18.2 ± 0.127	10	7.55 ± 0.66	13.2 ± 0.26
J.57	-1.25 ± 1.15	11	-0.653 ± 0.58	6.10 ± 0.37

^a Maximum SB decrement, values are in units of 10^{-18} ergs sec $^{-1}$ cm $^{-2}$ arcsec 2 .

^b Radius of SB $_1$ contour in kpc.

^c Measured from EW images.

^d Measured from Keck/LRIS spectra.

tection limit using the largest area is $SB(\text{Mg II}) = 5.74 \times 10^{-19}$ ergs sec $^{-1}$ cm $^{-2}$ arcsec 2 , computed for the galaxy J.03.

6. DISCUSSION

6.1. Previous Detections of Extended Mg II Emission

Previous constraints on the brightness of scattered Mg II emission were reported by Rubin et al. (2011). In this work the authors studied emission from the starburst galaxy TKRS 4389 at $z = 0.69$ with a star formation rate of $49.8 M_{\odot} \text{ yr}^{-1}$. This emission was detected in a 2-dimensional Keck/LRIS spectrum, with flux from the emission reaching (8.0 ± 0.4) and $(4.4 \pm 0.4) \times 10^{-18}$ ergs sec $^{-1}$ cm $^{-2}$ at λ_{2796} and (4.0 ± 0.3) and $(2.5 \pm 0.4) \times 10^{-18}$ ergs sec $^{-1}$ cm $^{-2}$ at λ_{2803} in two independent locations spatially offset from the galaxy continuum. The flux from the emission can be converted into two surface brightness values by taking the average of the flux measured at each location and each transition, and dividing by a 1 sq.arcsec aperture.

Figure 9 shows a plot of the 5σ detection limits determined for each galaxy in our sample as well as the SB calculated for the galaxy TKRS 4389 vs SFR. The figure suggests that we should be able to detect scattered Mg II emission with strengths similar to that detected in TKRS 4389. Taken at face value the figure could be consistent with a positive correlation between Mg II SB and SFR. Future observations are needed to verify this trend, and thus the possibility of detect Mg II in emission for high SFR, $> \sim 50 M_{\odot} \text{ yr}^{-1}$. Figure 10 shows the same SB values as Figure 9 vs $\log(M_*/M_{\odot})$. The figure is suggestive that Mg II is stronger in emission in lower mass galaxies, and is consistent with the results in Erb et al. (2012) and Feltre et al. (2018).

6.2. Geometry of Scattering Material

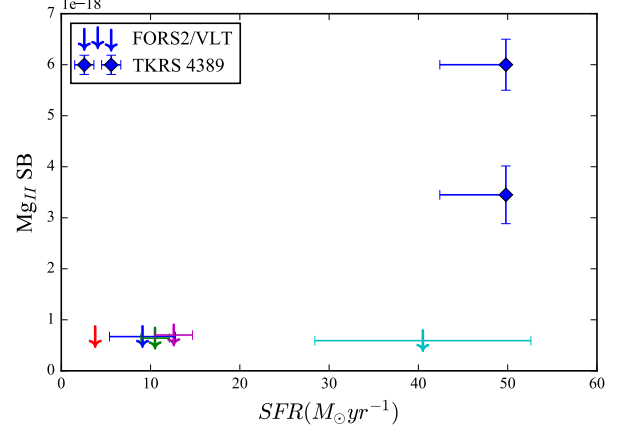


Figure 9. Comparison of our detection limits to the measured extended emission of TKRS 4389 vs SFR. The arrow color-galaxy scheme is: (Blue) J.26, (Green) J.36, (Red) J.03, (Cyan) J.64, (Magenta) J.57. Our imaging is sufficiently sensitive to detect extended emission at similar strengths to the extended emission measured for TKRS 4389 with $SFR \sim 50 M_{\odot} \text{ yr}^{-1}$.

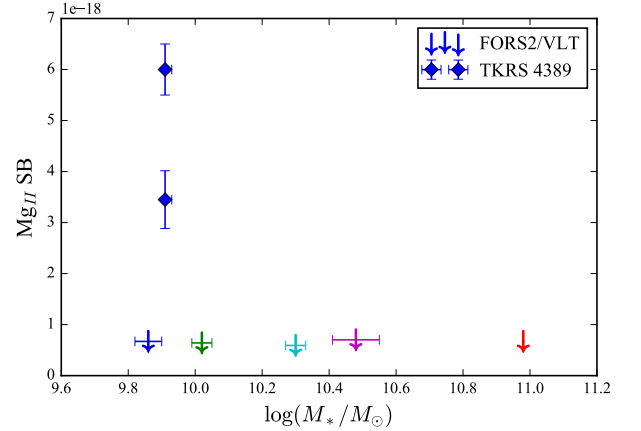


Figure 10. Comparison of our detection limits to the measured extended emission of TKRS 4389 vs Mass. Our imaging is sufficiently sensitive to detect extended emission at similar strengths to the extended emission measured for TKRS 4389 with mass $\sim 10 \log(M_*/M_{\odot})$. The color scheme is the same as Figure 9.

In idealized wind models of cool gas outflows, discussed in Prochaska et al. (2011), radiative transfer calculations suggest that strong Mg II emission is predicted to be generated along with ubiquitous blueshifted absorption of Mg II. For isotropic and dust-free scenarios, photons are conserved, as any absorbed continuum photon is eventually re-emitted, most likely at a different rest wavelength. Therefore, the total equivalent width of both the absorption and emission features is equal to

zero in such models. Assuming that our galaxies host an isotropic and dust-free wind, we wish to determine how much emission is predicted to be generated by this wind, and how the SB of this emission compares to our detection limits.

To calculate the predicted emission flux we first determine the flux absorbed by Mg II ions. Using our Keck/LRIS spectra, we find the average value of the continuum near the Mg II doublet and multiply this value by the observed EW of the doublet. Then to estimate the SB, we distribute this flux uniformly inside multiple annuli of varying sizes. These annuli all have an inner radius equal to the galaxy’s isophotal radius and successively larger outer radii. Additionally, since our SB limits are dependent on the size of the aperture used, we calculate the SB detection limits of our images inside each of the aforementioned annuli. Figure 13 shows how the SB of emission varies with the spatial extent of the annulus (black squares), as well as how the SB compares with our detection limits (thin black curve). Excepting the galaxy J.03, the SB of this “missing” emission lies above our detection limits. Under the assumption that the wind in these galaxies does in fact extend beyond the SB₁ isophotal contour (at $R_{\perp} = 8 - 21$ kpc), this result suggests that these galaxies do not host isotropic, dust-free winds.

6.2.1. Anisotropic and Dust Free Winds

There are many phenomena that may reduce the SB of the scattered Mg II emission so that it is consistent with our observations. One factor that can affect the observed emission strength is morphology of the wind. Anisotropic winds were shown in Prochaska et al. (2011) to exhibit reduced emission strength compared to an isotropic wind. Direct evidence of anisotropic winds, specifically bipolar morphology, has been observed in emission from cold and shock heated gas in local starburst galaxies (e.g. Walter et al. 2002; Westmoquette et al. 2008; Strickland & Heckman 2009). For distant galaxies, enhanced absorption Mg II along a galaxy’s minor axis (Bordoloi et al. 2011; Kacprzak et al. 2012; Bouché et al. 2012), higher outflow velocities toward $z \sim 1$ galaxies having $i < 45^\circ$ are suggestive of bipolar outflows. Furthermore, Rubin et al. (2014) analysis provide strong evidence for the ubiquity of bipolar outflows in distant galaxies.

We now assume that the emission in our galaxies is reduced by the effect of anisotropy. For the anisotropic winds modeled in Prochaska et al. (2011), the emission is reduced by the factor $\Omega/4\pi$, where Ω is the angular extent of the wind. As Prochaska et al. (2011) points out, given that the outflow must cover most of the con-

tinuum in order to be detected, the value of Ω has an approximate lower limit of $\Omega > 2\pi$. We show the predicted SB profiles for wind emission from our galaxies assuming $\Omega = 2\pi$ with blue diamonds in Figure 13.

After reducing the SB of the expected Mg II emission by the corresponding factor of 2, we predict profiles that fall below our SB detection limits for galaxies J.26 and J.36. However, the SB profile of J.64 remains above our SB detection limits, suggesting additional phenomena are needed to reduce the strength of scattered emission. Previously discussed in Section 5, this object is the intrinsically brightest galaxy of our sample and exhibits the strong Mg II absorption, which suggests a strong ISM component. Prochaska et al. (2011) note that Mg II photons can be more effectively trapped in objects with a large amount of dusty interstellar material.

6.2.2. Anisotropic and Dusty Winds

Dust in the wind is another factor that can reduce the observed emission strength, and affect the shape of the line profile. In the Prochaska et al. (2011) radiative models that include dust in the wind material, the dominant effect is that the most redshifted emission is suppressed. The strength of the flux is reduced by a factor of $(1 + \tau_{\text{dust}})^{-1}$, where τ_{dust} is the integrated opacity of dust. Estimates of the dust content of each galaxy come from SED modeling in Rubin et al. (2014) and are shown in Table 1. The green hexagons in Figure 13 shows the predicted SB profile for an anisotropic wind with dust. For the galaxies J.26, J.36 and J.03, the introduction of dust reduces the predicted emission further below our detection limits and profiles reduced by anisotropy only. For galaxy J.64, anisotropy alone does not reduce the predicted emission of J.64 below our detection limits. The profile in Figure 13 shows that combination of dust and anisotropy is sufficient in reducing the strength of the scattered emission of J.64 below our detection limits.

7. CONCLUSION

We presented the results from a narrowband imaging search for Mg II emission in a sample of star-forming galaxies at a redshift of $z \sim 0.70$ which are known to drive outflows in Mg II. We did not detect any Mg II emission in this sample, and place a 5σ detection upper limit of $\text{SB}(\text{Mg II}) \approx 5.74 \times 10^{-19} \text{ ergs sec}^{-1} \text{ cm}^{-2} \text{ arcsec}^2$, measured in an extended annulus, $\sim 20 \text{ sq.arcsec}$, at minimum projected distances of 13, 18, 21 and 24 kpc from the centers of each target. We detected a significant amount of Mg II absorption in a total of 4 galaxies out of our 5 galaxy sample. Furthermore, our imaging allows us to generate spatially-resolved maps of Mg II absorption in a distant galaxy sample for the first time. This absorption covers the center of the galaxies

out to an isophotal radius defined by the SB_1 (1σ) depth of a continuum + Mg II image, at approximately ~ 22 kpc, suggesting that the absorbing gas is optically thick and fully covers the stellar disks out to this distance. The EWs measured in these maps are in broad agreement with measured EWs using Keck/LRIS integrated slit spectroscopy. Additionally, our radial projections of the mean EW for our sample galaxies suggest that the EWs due to Mg II are approximately constant across the galaxies surfaces.

We compared our detection limits with the prediction of the radiative transfer models of Prochaska et al. (2011). We are able to rule out that the winds in our

sample are isotropic and dust free, as our images are sufficiently sensitive to detect the predicted emission from such models. By assuming dust/dust-free anisotropic winds, we were able to reduce the strength of the predicted Mg II emission to lie below our detection limit for 4 of our sample galaxies. Although our detection limits suggest that the winds in our sample are not isotropic and dust-free, questions linger regarding the relative roles wind anisotropy, dust content, and extent play in reducing scattered emission. Thus, Deeper imaging in Mg II are required to observe any extended emission beyond the depth of our images and up to our predicted SB profiles to characterize the morphology of these winds.

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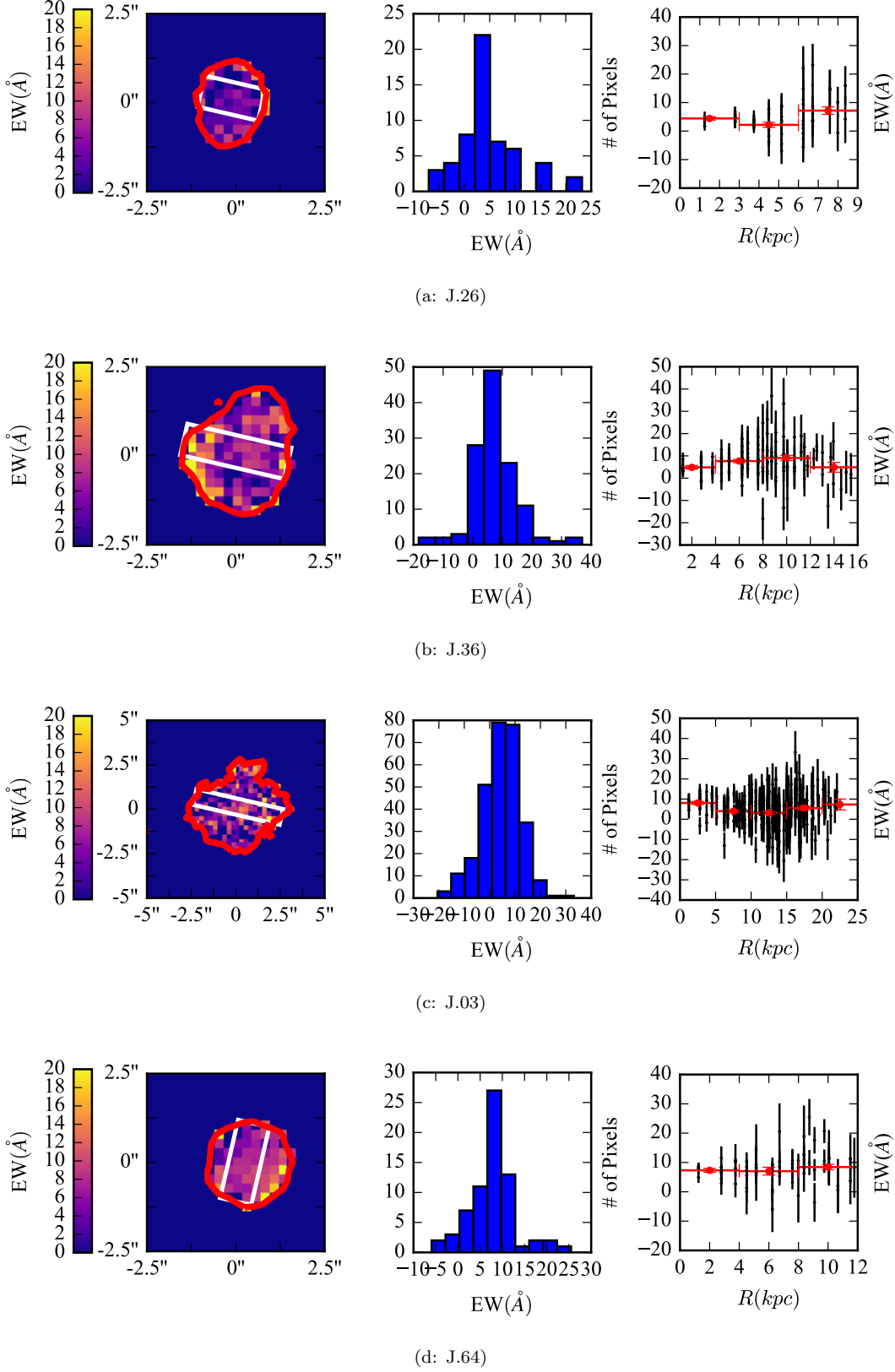


Figure 11. (Left) Stamp of the equivalent widths (EWs) inside the red 1SB₁ contour. The white contour represents the 0.9'' aperture slit used to measure the equivalent width of absorption in the Keck/LRIS spectra. (Middle) Distribution of absorption pixels with continuum flux S/N greater than 1.5 inside the slit aperture. (Right) The EWs of absorption projected in distance from center are in blue and binned EW measurements red. The error bars in the radial axis represent the width of the radial bin used.

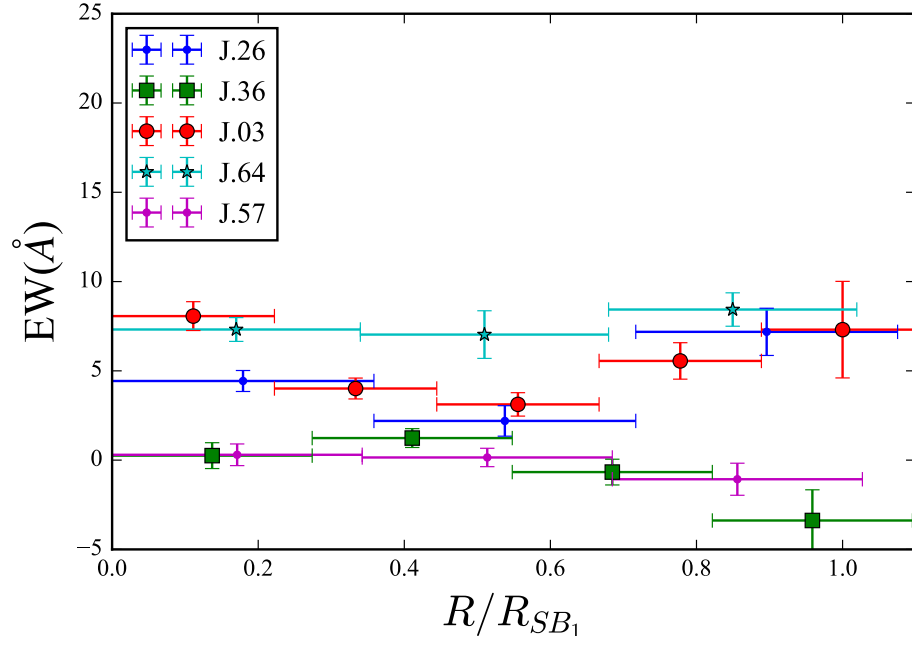


Figure 12. Normalized radial profile of the mean EW of Mg II absorption for all galaxies. Measurements for the mean EW binned in radial increments compiled to test the significance of any trends found in the individual galaxy profiles. There is significant scatter relative to the error bars at extended distances.

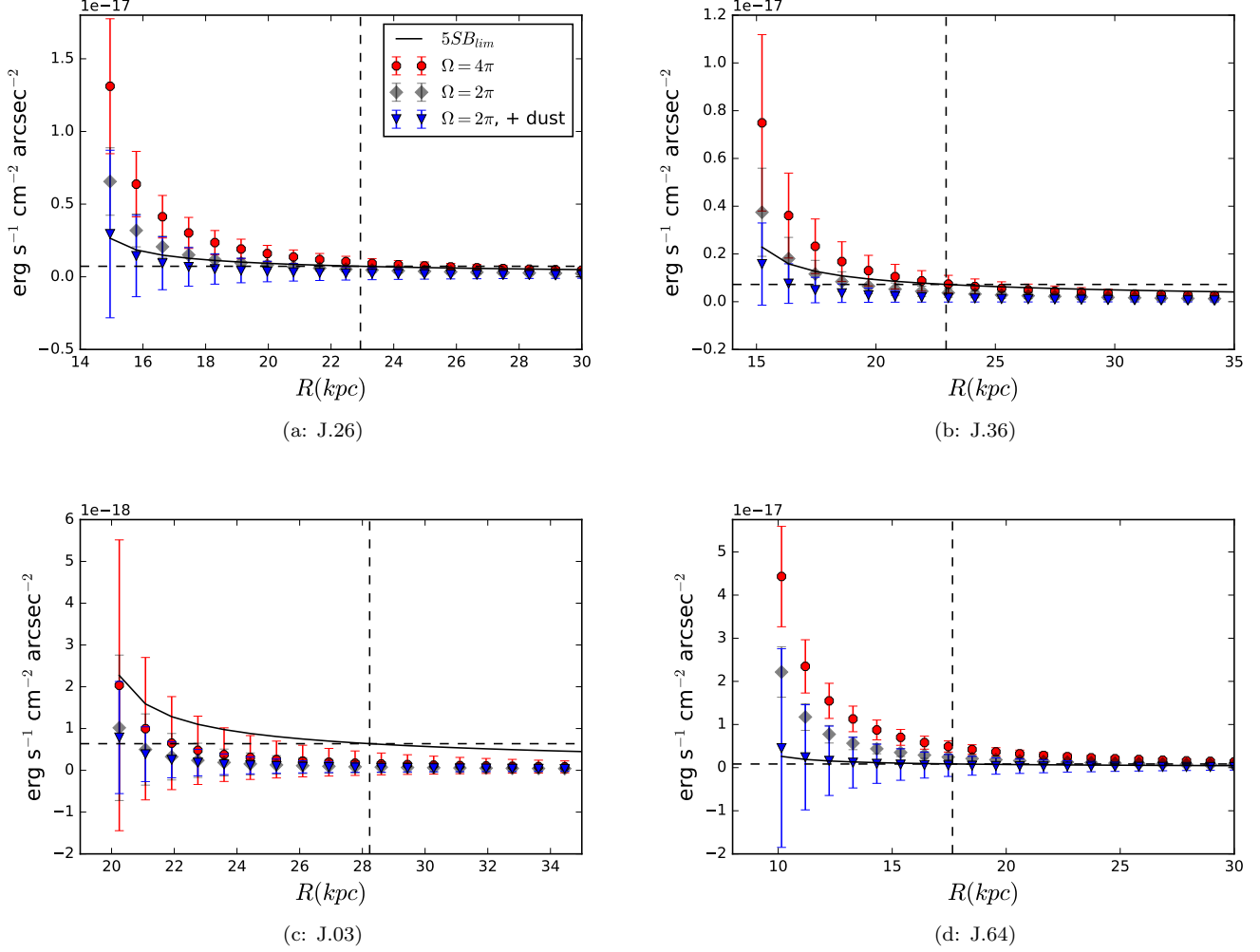


Figure 13. Estimated SB profile of Mg II emission for isotropic and anisotropic dust-free winds. The red hexagons points are SB of predicted Mg II emission uniformly distributed inside an annulus for an angle $\Omega = 4\pi$, or an isotropic wind. The black diamond points are SB of missing Mg II emission uniformly distributed inside an annulus for an angle $\Omega = 2\pi$, or an anisotropic wind. The blue triangle points are SB of predicted emission for a wind that is anisotropic and affected by dust. The solid line shows the value of the SB detection limits. The dashed vertical and horizontal lines represent outer radius of the extended annulus used to measure the initial detection limits and the value of the SB detection limit in that annulus. The legend in panel (a) holds for the remaining panels.