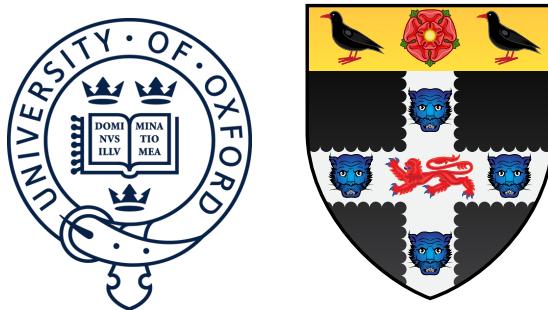


# First Measurement of the Solar Neutrino Oscillation Parameters via Boron-8 Solar Neutrinos in SNO+



**Daniel Cookman**

Christ Church College

University of Oxford

A thesis submitted in fulfilment of the requirements for the degree of

*Doctor of Philosophy*

Trinity Term 2022



*To my parents*

*and*

*To Oscar Jacobsson:*

*The best of us*



## **Acknowledgements**

And I would like to acknowledge ...



## **Abstract**

Formal 1-page summary of the work completed in the thesis.



# Table of contents

<b>List of Figures</b>	<b>xiii</b>
<b>List of Tables</b>	<b>xv</b>
<b>List of Acronyms</b>	<b>xvii</b>
<b>Introduction</b>	<b>1</b>
<b>1 The Theory of Neutrino Physics</b>	<b>3</b>
1.1 The Standard Model and Neutrinos . . . . .	3
1.1.1 A Brief Introduction to the Standard Model . . . . .	3
1.1.2 Neutrinos within the Standard Model . . . . .	4
1.2 Neutrino Oscillations and Neutrino Masses . . . . .	5
1.2.1 The Evidence for Neutrino Oscillations . . . . .	5
1.2.2 The Phenomenology of Neutrino Oscillations . . . . .	6
1.2.3 The Origins of Neutrino Mass . . . . .	6
<b>2 The SNO+ Detector</b>	<b>9</b>
2.1 Detector Geometry . . . . .	9
2.2 Experimental Phases . . . . .	10
2.3 Detecting and Recording an Event in SNO+: A Journey . . . . .	12
2.3.1 Particle Interactions with Matter . . . . .	12

2.3.2	Optical Processes . . . . .	17
2.3.3	Detection by PMTs . . . . .	24
2.3.4	Data Acquisition and Triggering . . . . .	27
2.3.5	Operation of the Detector . . . . .	29
2.4	Calibrations and Detector Modelling . . . . .	30
2.4.1	Detector Monitoring . . . . .	30
2.4.2	Electronic and PMT Calibrations . . . . .	31
2.4.3	Energy and Optical Calibrations . . . . .	32
2.4.4	Event Reconstruction . . . . .	34
2.4.5	Event Simulation . . . . .	37
<b>3</b>	<b>The SMELLIE Calibration System</b>	<b>39</b>
3.1	The SMELLIE Hardware . . . . .	40
3.2	Software for SMELLIE Data-taking . . . . .	40
3.3	Commissioning SMELLIE in the Scintillator Phase . . . . .	41
<b>4</b>	<b>Simulating SMELLIE Events</b>	<b>43</b>
4.1	Improving the SMELLIE Generator Algorithm . . . . .	44
4.1.1	Previous Attempts at SMELLIE Event Simulation . . . . .	44
4.1.2	The new generator . . . . .	47
4.2	Improving the beam profiles . . . . .	51
4.2.1	Combining beam profile datasets . . . . .	54
4.2.2	Results & Discussion . . . . .	60
4.3	Comparisons between Data and Simulation . . . . .	64
4.3.1	Forward Hemisphere Discrepancies . . . . .	65
4.3.2	Emission Time Discrepancies . . . . .	65
4.3.3	Backward Hemisphere Discrepancies . . . . .	65

---

Table of contents	xi
<b>5 Analysis of SMELLIE Data in the Scintillator Phase</b>	<b>67</b>
5.1 Extinction Length Analysis . . . . .	67
5.1.1 Motivation . . . . .	67
5.1.2 Analysis Approach . . . . .	68
5.1.3 Validation of the Analysis in Simulation . . . . .	69
5.1.4 Results in Data . . . . .	69
5.2 Scattering Analysis . . . . .	70
5.2.1 Historical Approaches and the Problem of Systematics . . . . .	70
5.2.2 New Methodology . . . . .	71
5.2.3 Results . . . . .	72
<b>6 Solar Oscillation Analysis</b>	<b>73</b>
6.1 Analysis Methodology . . . . .	74
6.1.1 Observational Principle . . . . .	74
6.1.2 Background Processes . . . . .	77
6.1.3 The Log-likelihood Test Statistic . . . . .	83
6.1.4 The Bayesian Statistical Approach . . . . .	85
6.1.5 Markov Chain Monte Carlo . . . . .	87
6.1.6 Choosing Priors . . . . .	89
6.1.7 Including Systematics in the Fit . . . . .	89
6.1.8 Including Oscillations in the Fit . . . . .	91
6.2 Analysis on Scintillator-Phase data . . . . .	93
6.2.1 Dataset and Livetime . . . . .	93
6.2.2 Event Selection . . . . .	95
6.2.3 Expected Rates and Constraints . . . . .	98
6.2.4 Results . . . . .	98
6.3 Sensitivity Projections . . . . .	100

<b>7 Conclusions and Suggestions for Future Work</b>	<b>101</b>
<b>References</b>	<b>105</b>

# List of Figures

2.1	3D model of the SNO+ detector . . . . .	10
2.2	The current modelled distribution for the emission timing of the SNO+ scintillator . . . . .	15
2.3	Comparison of the SNO+ liquid scintillator cocktail's absorption and emission properties, with/without BisMSB. . . . .	17
2.4	Measured properties of the UPW and acrylic in the water phase. . . . .	22
2.5	Refractive indices of acrylic, UPW, and LABPPO as a function of wavelength . . . . .	23
2.7	Efficiencies of the R1408-type PMTs used as standard within SNO+ . .	26
2.8	Example charge spectra for a PMT as a function of the true npe generated	26
2.9	Schematic of the front-end electronics used for data acquisition and triggering in SNO+ . . . . .	27
4.1	Comparison of SMELLIE between data and MC using a 1D generator .	45
4.2	Number of attempts taken for existing 2D SMELLIE generator to accept direction . . . . .	48
4.3	Intensity map of existing beam profile, and interpolated sampling map	49
4.4	Histogram of interpolated intensities within the 2D direction-space . .	50

4.5 Comparison of water-phase data to MC generated using both the old and new 2D beam profile generator approaches, with the updated beam profiles . . . . .	52
4.6 Plot of $\Delta\chi^2$ for both single subruns of a typical PMT, and when all relevant subruns are combined . . . . .	59
4.7 Comparison between old and updated beam profiles for fibre FS055, after combining multiple data sets . . . . .	61
4.8 Interpolated intensity map for the new updated beam profile of fibre FS055 . . . . .	62
4.9 Residuals from subruns at two different wavelengths, both compared to the combined beam profile model for fibre FS055 . . . . .	64
6.1 True ${}^8\text{B}$ $E_\nu$ versus true $E_e$ . . . . .	75
6.2 The evolution of energy distributions related to ${}^8\text{B}$ solar neutrino detection. . . . .	76
6.4 The ${}^{238}\text{U}$ and ${}^{232}\text{Th}$ decay chains. . . . .	78
6.6 Schematic of $(\alpha, n)$ interactions. . . . .	80

# List of Tables

2.1	Current values used to model scintillator emission from electrons in 2.2 g/L LABPPO. . . . .	15
4.1	Water-phase runs used for new beam profiling. . . . .	55
6.1	Cuts used in this analysis. . . . .	95
6.2	Impact of each cut on the quantity of events in data. . . . .	97
6.3	Impact of each cut on the quantity of events in data. . . . .	97



# List of Acronyms

**AmBe** Americium-Beryllium radioactive source

**AMELLIE** Attenuation Module for ELLIE

**AV** Acrylic vessel

**BD** 1,2-Butanediol

**BHT** Butylated hydroxytoluene

**BisMSB** 1,4-Bis(2-methylstyryl)benzene

**CTC** Crate Trigger Card

**DAQ** Data acquisition (system)

**DB** Daughter Board

**DDA** N,N-Dimethyldodecylamine

**ECA** Electronic Calibration

**ELLIE** Embedded LED/Laser Light Injection Entity

**EXTA** External Asynchronous (Trigger)

**FEC** Front-End Card

**GTID** Global Trigger Identification number

**GT** Global Trigger

**IW** In-window events

**LAB** Linear alkylbenzene

**MTC/A+** Analogue Master Trigger Card

**MTC/D** Digital Master Trigger Card

**npe** Number of photoelectrons

**OOW** Out-of-window events

**OWLs** Outward-looking PMTs

**PCA** PMT Calibration

**PMTIC** PMT Interface Card

**PMT** Photomultiplier Tube

**PPO** 2,5-Diphenyloxazole

**PSUP** PMT support structure

**QE** Quantum efficiency (of a PMT)

**QHL** Charge with high gain over a ‘long’ integration time (390 ns)

**QHS** Charge with high gain over a ‘short’ integration time (60 ns)

**QLX** Charge with low gain over a ‘long’ integration time (390 ns)

**SMELLIE** Scattering Module for ELLIE

**SNO** Sudbury Neutrino Observatory

**TAC** Time-to-amplitude Converter

**TeA** Telluric acid, Te(OH)<sub>6</sub>

**TELLIE** Timing subsystem for ELLIE

**TeLS** Tellurium-loaded liquid scintillator

**TOF** Time-of-flight (of a photon)

**TTS** Transit time spread (of a PMT)

**TUBii** Trigger Utility Board Mark ii

**UPW** Ultra-pure water

**ZDAB** Zebra Database (file format)

**nhit** Number of PMT hits in an event

**RATDB** RAT Database

**RAT** Reactor Analysis Tool

**ADC** Analogue-to-Digital Converter



# Introduction

Couple of pages outlining document's structure and contents (this is what each of the chapters is here for). Less formal than the abstract, also explaining what the expected audience of this thesis is: who will find this document useful!

1

2

3

4



# Chapter 1

## The Theory of Neutrino Physics

*Light*

*Light*

*The visible reminder of Invisible Light*

---

*The Rock*

T. S. ELIOT

### 1.1 The Standard Model and Neutrinos

#### 1.1.1 A Brief Introduction to the Standard Model

Covering how the SM works at the highest level, including:

- Quantum Field Theory and the Lagrangian dynamical framework

- The connection between symmetries of a QFT model and its gauge fields that describe the model's forces

- The SM's fundamental symmetries, and associated forces, but —

- 1     • Not (exactly) what we see “normally”! The electromagnetic and weak forces  
2       appear distinct, and the weak gauge bosons have mass. To explain this, we need  
3       a further component, the Brout-Englert-Higgs (BEH) Mechanism.
- 4 [2 pages total]

5 **1.1.2 Neutrinos within the Standard Model**

- 6     • Basic description of where neutrinos fit into SM: 3 kinds of neutral fermion, the  
7       counterparts to the charged fermions. Interacts with the weak force only.
- 8     • Summary of the experimental evidence for this picture: mainly, the discovery of  
9       electron anti-neutrinos by Cowan and Reines, the muon neutrino by Lederman,  
10       Schwartz, and Steinberger, and the tau neutrino by the DONUT Collaboration.  
11       Further critical experiments include the first measurement of a neutrino’s helicity  
12       by Goldhaber et al. as well as Danby et al.’s demonstration that  $\nu_\mu$  are distinct  
13       from  $\nu_e$ .
- 14     • More detailed description, via Feynman diagrams, of the two fundamental modes  
15       of interaction by neutrinos with the weak force: charged- and neutral-current  
16       interactions. A brief mention of the quantitative theory that underlies description:  
17       Gashow, Salam, and Weinberg’s Electroweak Theory. This explains not only the  
18       V–A structure of charged-current interactions, but also predicted accurately the  
19       nature of neutral-current interactions. (Given space constraints, I see no reason  
20       to go into much of the details of the theory, or the many experimental tests of  
21       its structure.)

22 [4 pages]

## 1.2 Neutrino Oscillations and Neutrino Masses

1

### 1.2.1 The Evidence for Neutrino Oscillations

2

- Describe status quo ante of massless nature of neutrinos: BEH mechanism as exists cannot allow for neutrinos to have mass as only left-handed neutrinos have been observed.  
3  
4  
5
- Furthermore, strong experimental limits on neutrino masses, from e.g. tritium-decay endpoint measurements by the KATRIN experiment and cosmological inferences from the CMB by the Planck satellite.  
6  
7  
8
- But — then neutrino oscillations are observed over a variety of experiments and contexts. Summarise critical bits of evidence:  
9  
10
- Electron neutrino disappearance in solar neutrino experiments, including Ray Davis' Homestake experiment, the SAGE/GALLEX experiments, and SNO. For the latter, the comparison of charged-current and neutral-current modes of interaction was clear evidence of neutrino oscillations over other types of process (e.g. neutrino decay).  
11  
12  
13  
14  
15
- Include in the above a brief description of Bahcall's Standard Solar Model.  
16
- Muon neutrino disappearance in atmospheric and long-baseline accelerator neutrino experiments, such as Super-Kamiokande, T2K, and No $\nu$ a.  
17  
18
- A few further observations to note are: reactor electron anti-neutrino disappearance from both KamLAND and Daya Bay; tau neutrino appearance at the OPERA experiment; short-baseline neutrino anomaly within LSND and MiniBooNE (with recent contrary evidence from MicroBooNE).  
19  
20  
21  
22

**1.2.2 The Phenomenology of Neutrino Oscillations**

- Describe the current phenomenological model of 3-flavour neutrino oscillations that can explain all of this evidence: the PMNS mixing matrix.
- Describe also the MSW effect, which is critical for explaining solar neutrino oscillations.
- Show the formula for solar neutrino oscillations, given this MSW effect in both the Sun and Earth. Note the dependence of solar neutrino oscillations on only the “solar” oscillation parameters. This is all particularly useful for the solar analysis chapter.

$$P_{ee} \left( \tan 2\theta_{12}^M, \sin \theta_{13}^M, \Delta m_{21,M}^2 \right) = BLAH \quad (1.1)$$

[3 pages]

**1.2.3 The Origins of Neutrino Mass**

- Observed neutrino oscillations require at least two neutrino mass states to be non-zero. Given constraints of the current SM, two main ways of adding neutrino masses: a Dirac mass term (i.e. allowing for sterile neutrinos), and a Majorana mass term.
- For latter, briefly describe what a Majorana particle is, and how with the Seesaw Mechanism (just the simple Type 1 described in-text) one can not only get neutrino masses but also explain their lightness relative to the other massive SM particles. Note that there exist more elaborate versions of this theory.
- Furthermore, with reference to the Sakharov conditions, describe qualitatively how the Seesaw Mechanism also allows for possible leptogenesis/baryogenesis in the early Universe, and hence could explain its matter-antimatter asymmetry.

---

1.2 Neutrino Oscillations and Neutrino Masses

## 7

- Describe briefly the nuclear physics behind double-beta decay (i.e. why it can happen at all over just normal beta decay), and then how Majorana neutrinos allow for neutrinoless double beta decay,  $0\nu\beta\beta$ .  
1  
2  
3
- Describe the experimental signature of  $0\nu\beta\beta$ : a spike of events of observed energy equal to the Q-value of the decay.  
4  
5
- Note Schecter-Valle Theorem ensures that any observation of  $0\nu\beta\beta$  must be the result of neutrinos being Majorana. I.e. the Universe cannot conspire against us and have  $0\nu\beta\beta$  without Majorana neutrinos.  
6  
7  
8
- Very briefly note the current status of the search for  $0\nu\beta\beta$ , describing the main varieties of experimental setup seen, along with a nice canonical example of such an experiment and their best limit. In particular, the Germanium-crystal detectors such as GERDA, Xenon-TPC detectors like EXO-200, and large-scale liquid scintillators such as KamLAND-Zen.  
9  
10  
11  
12  
13

[3 pages]

14

[CHAPTER TOTAL: 17 pages]

15



# Chapter 2

## The SNO+ Detector

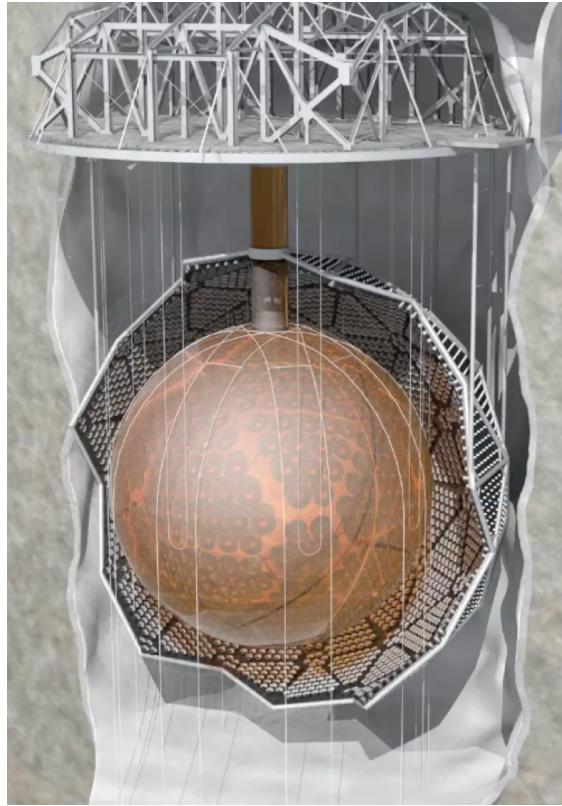
*The light-soaked days are coming.*

---

JOHN GREEN

### 2.1 Detector Geometry

The SNO+ detector is a large, multi-purpose neutrino detector built in the SNOLAB underground laboratory near Sudbury, Canada. Its main detector structure is taken from the Sudbury Neutrino Observatory (SNO) [1], which can be seen in Fig. 2.1. At the heart of the detector lies the main detector medium, which changes depending on the phase of the experiment — more on the specifics of this in Section 2.2. This medium is held within a 12 m diameter sphere known as the Acrylic Vessel (AV). The AV floats within a body of ultra-pure water (UPW), beyond which is a stainless steel support structure (PSUP) that holds 9362 inward-facing Photomultiplier Tubes (PMTs). It is these PMTs that detect the light generated from physics events that occur within the detector medium. The AV is kept in place relative to the PSUP through a series of ‘hold-up’ and ‘hold-down’ tensylon ropes. All of these components are suspended within a large cylindrical cavity also filled with UPW. 91 outward-looking PMTs (OWLs) are



**Fig. 2.1:** 3D model of the SNO+ detector [2].

<sup>1</sup> also affixed to the outside of the PSUP, allowing for the effective vetoing of cosmic ray  
<sup>2</sup> muons.

<sup>3</sup> Directly above the detector is the Deck, within which all the detector electronics  
<sup>4</sup> are kept. Access within the AV for calibration tools and filling is possible only through  
<sup>5</sup> the acrylic ‘neck’ on top of the AV. Full details of the design of the current detector  
<sup>6</sup> can be found in [2].

## <sup>7</sup> 2.2 Experimental Phases

<sup>8</sup> As mentioned earlier, SNO+ was designed to fulfil a number of physics goals over  
<sup>9</sup> multiple ‘phases’ of the detector’s lifetime. The phases are distinguished by the medium  
<sup>10</sup> that fills the AV. The first main phase (after a brief **Air Fill Phase** used only for

---

**2.2 Experimental Phases****11**

detector commissioning) was that of the **Water Fill Phase**, with data taken between 1 May 2017 and July 2019. This was used to perform fundamental optical calibrations 2 of the detector [3], measurements of the solar neutrino flux [4], observation of neutrino 3 oscillations in reactor anti-neutrinos [5], and searches for nucleon decay [6, 7]. 4

After this, the detector was filled with 800 tonnes of liquid scintillator known as 5 linear alkylbenzene (LAB), mixed with the fluor 2,5-diphenyloxazole (PPO). More 6 information on the physics of scintillators can be found in Section 2.3.1. Filling of the 7 LABPPO cocktail had to be paused in March 2020 due to the COVID-19 pandemic, 8 leading to the detector having its bottom half still filled with UPW, and the top 9 half filled with LAB and PPO at 0.5 g/L. This impromptu phase became known as 10 the **Partial Fill**, and allowed for some creative analyses to be performed: an initial 11 neutrino oscillation analysis from reactor anti-neutrinos [8], as well as the first ever 12 observation of directionality in a high light yield scintillator [9, 10]. Eventually, filling 13 of the detector with liquid scintillator completed in May 2021. At that point, the 14 concentration of PPO in the detector was at 0.6 g/L, markedly below the target level 15 of 2.0 g/L. A further ‘PPO top-up’ campaign then proceeded, finishing in April 2022 16 with a final concentration of 2.2 g/L PPO. Thus began the **Scintillator Phase** of 17 the experiment, which continues on during the time of writing. The main goals for 18 this phase include a number of solar neutrino analyses (including the one described in 19 Chapter 6), a precision measurement of the neutrino oscillation parameter  $\Delta m_{21}^2$  using 20 reactor anti-neutrinos [8], and further calibrations of the detector and its backgrounds. 21

Finally, in the near future the detector will be loaded with Tellurium for the 22 **Tellurium Phase**, allowing for the flagship analysis of the experiment to begin: 23 neutrinoless double beta decay. In order to load Te within the liquid scintillator 24 in a stable manner, a chemical loading process has been developed, as described 25 in [11]. The Te starts within  $\text{Te(OH)}_6$  (telluric acid, otherwise known as TeA), which 26

<sup>1</sup> after purification will be reacted with 1,2-butanediol (BD) via heating and addition  
<sup>2</sup> of N,N-Dimethyldodecylamine (DDA), which acts as a stabiliser. What results is  
<sup>3</sup> tellurium-loaded scintillator, TeLS.

<sup>4</sup> Two further chemicals are planned to be added to the scintillator cocktail. The  
<sup>5</sup> antioxidant butylated hydroxytoluene (BHT) will be added to capture any free-radicals  
<sup>6</sup> within the liquid scintillator, hopefully preventing any oxidation reactions that could  
<sup>7</sup> lead to the ‘yellowing’ of the scintillator, a degradation of its optical properties. The  
<sup>8</sup> addition of BHT is not expected to impact the detector’s optics in any substantial  
<sup>9</sup> way. However, the other substance to also be added, 1,4-Bis(2-methylstyryl)benzene  
<sup>10</sup> (bisMSB), will impact the optics. BisMSB is a ‘wavelength-shifter’ which enables the  
<sup>11</sup> scintillator cocktail to generate light with a greater overall detection efficiency — more  
<sup>12</sup> on the details of this in Section 2.3.1.

<sup>13</sup> **2.3 Detecting and Recording an Event in SNO+:**  
<sup>14</sup> **A Journey**

<sup>15</sup> To understand the SNO+ detector well, it is worth thinking about how the information  
<sup>16</sup> contained in a physics event, e.g. a solar neutrino interaction, gets observed. This  
<sup>17</sup> section follows the journey of such an event.

<sup>18</sup> **2.3.1 Particle Interactions with Matter**

<sup>19</sup> All observable physics events within the detector begin by the generation of some  
<sup>20</sup> form of ionising radiation:  $\alpha$ ,  $\beta^\pm$ ,  $\gamma$ ,  $p$  or  $n$ . These can be created via numerous  
<sup>21</sup> processes, both exciting (e.g.  $0\nu\beta\beta$  or interactions of neutrinos) and annoying (e.g.  
<sup>22</sup> decay of background radioisotopes): see Section 6.1.2 for some of them. Regardless of  
<sup>23</sup> their origin, these particles begin propagating through the detector, and interacting

## 2.3 Detecting and Recording an Event in SNO+: A Journey

13

with the detector medium. A number of mechanisms then allow for the generation of optical-wavelength light as a result of these interactions.

**Cherenkov Light Emission**

Whenever a charged particle travels through a dielectric medium at speeds faster than the speed of light in that medium, light is generated from the ‘wake’ of induced dipoles. This is known as **Cherenkov light**, a process much akin to the ‘sonic boom’ that occurs when an object travels at supersonic speeds through a medium. This light emanates outwards in a cone along the direction of the charge’s travel; the angle of the cone  $\theta_\gamma$  being purely a function of the speed of the charged particle relative to the speed of light in vacuum,  $\beta$ , and the refractive index of the medium  $n(\omega)$  at a given frequency  $\omega$ :  $\cos \theta_\gamma(\omega) = \frac{1}{n(\omega)\beta}$ . There is then a minimum speed necessary for Cherenkov light to be generated:  $\beta_{\min}(\omega) = 1/n(\omega)$ .

In addition to the characteristic cone shape of the light, the spectrum of the light generated is also distinctive. Igor Tamm and Ilya Frank determined the expected energy  $dE$  emitted per unit length travelled by the charged particle,  $dx$ , as [12]:

$$\frac{dE}{dx} = \frac{q^2}{4\pi} \int_{n(\omega)>1/\beta} \mu(\omega) \omega \left(1 - \frac{1}{\beta^2 n^2(\omega)}\right) d\omega. \quad (2.1)$$

Here,  $q$  is the charge of the moving particle and  $\mu(\omega)$  is the permeability of the medium.

All SNO+ detection media allow Cherenkov light to be generated, as long as sufficiently high energy particles traverse it. In the water fill phase of the detector, Cherenkov light was the only means by which light could be generated. Light from Cherenkov emission can still be created in liquid scintillator, but it tends to be swamped by another form of light generation: scintillation.

**1 Scintillation**

2 For certain special classes of material, the excitation and ionisation of atomic electrons  
 3 nearby a moving charged particle can lead to the generation of optical-wavelength  
 4 light, in a process known as **scintillation** (often generally referred to as ‘luminescence’  
 5 or ‘fluorescence’). Although multiple varieties of scintillator exist, the one used in  
 6 SNO+ is that of an organic liquid scintillator. For such liquids, scintillation light is  
 7 generated from the de-excitation of delocalised ‘ $\pi$ -electrons’ that are found within  
 8 certain carbon–carbon bonds [13]. A major example of these  $\pi$ -bonds are found in  
 9 benzene rings, which are present in LAB, PPO, and bisMSB.

10 Because of this delocalised structure, excited atomic  $\pi$ -electrons can stay in what  
 11 is typically the first-excited state for somewhat longer than typical excited states:  
 12 lifetimes of  $\mathcal{O}(10^{-9}$  s) as opposed to  $\mathcal{O}(10^{-12}$  s). This is what gives scintillation light  
 13 its characteristic ‘slow’ response relative to the instantaneous light generated by the  
 14 Cherenkov process. Moreover, decays from this state can emit light typically in the  
 15 optical-wavelength range. In addition to excited electrons, ionised electrons can also  
 16 recombine — that is, re-enter atomic orbitals — into various excited states, and then  
 17 decay back to the ground state, also allowing for the possibility of scintillation light to  
 18 be generated.

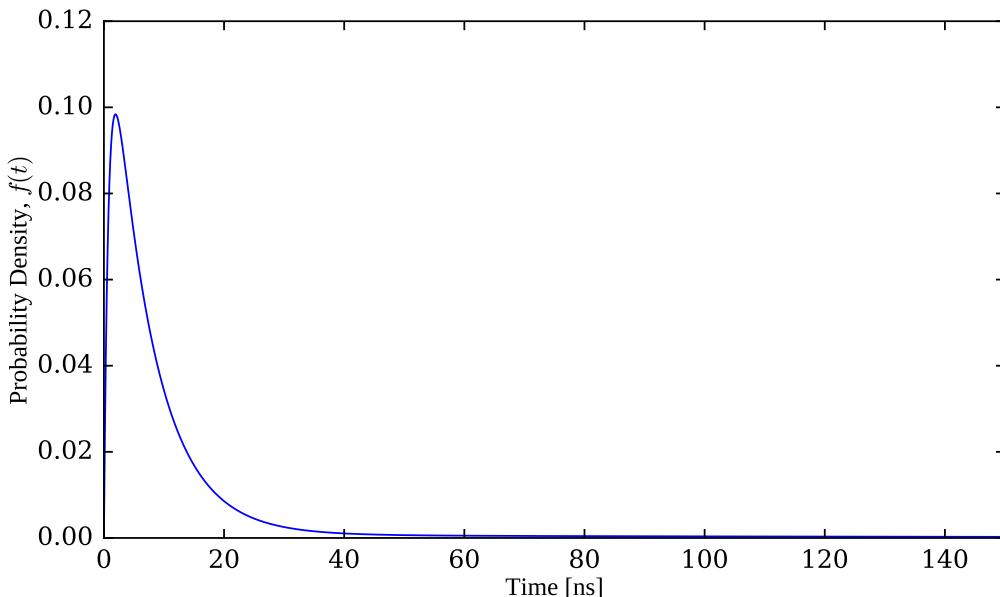
19 Because of atomic spin selection rules [13], scintillation light typically has, at the  
 20 very least, a ‘fast’ and ‘slow’ time component. In SNO+, we currently model emission of  
 21 scintillation light from LAB with 4 time components, following the timing distribution  
 22  $f(t)$  given by:

$$23 \quad f(t) = \sum_i A_i \left( \frac{e^{-t/\tau_i} - e^{-t/\tau_{\text{rise}}}}{\tau_i - \tau_{\text{rise}}} \right), \quad t > 0. \quad (2.2)$$

24 Here,  $A_i$  and  $\tau_i$  correspond to the fraction of light emitted and decay constant for each  
 25 component respectively, and  $\tau_{\text{rise}}$  is a common rise time. The current fitted values for

Component	$A_i$	$\tau_i$ [ns]
1	0.665	7.35
2	0.218	5.45
3	0.083	117.5
4	0.0346	425
Rise	–	0.8

**Table 2.1:** Current values used to model scintillator emission from electrons in 2.2 g/L LABPPO [14, 15].



**Fig. 2.2:** The current modelled distribution for the emission timing of the SNO+ scintillator, for electron tracks [14, 15].

these parameters for the emission from electron tracks can be seen in Table 2.1. A plot of  $f(t)$  with these values can be seen in Fig. 2.2.

When using just a single scintillating compound, the very same energy levels that can generate scintillation light are those that can absorb it. This can be a problem for large-scale detectors like SNO+, which depend on scintillation light being unobstructed in its path to the PMTs. Conveniently, this problem can be addressed with the addition of another scintillating component, known (somewhat confusingly) as the primary fluor. In SNO+, this is the PPO added to the LAB.

<sup>1</sup> When an LAB molecule is excited, that energy can be transferred to a PPO molecule  
<sup>2</sup> through what is known as a ‘non-radiative transfer’. In short, this transfer of energy  
<sup>3</sup> occurs not through the emission and absorption of optical photons, but through the  
<sup>4</sup> coupling of the molecules’ electric dipoles.<sup>1</sup> When the now-excited PPO molecule  
<sup>5</sup> de-excites to emit scintillation light, the different molecular structure it has generates  
<sup>6</sup> a different emission spectrum to that of LAB. These longer wavelengths of light are  
<sup>7</sup> no longer able to be absorbed by the LAB, allowing for a scintillator with less optical  
<sup>8</sup> absorption.

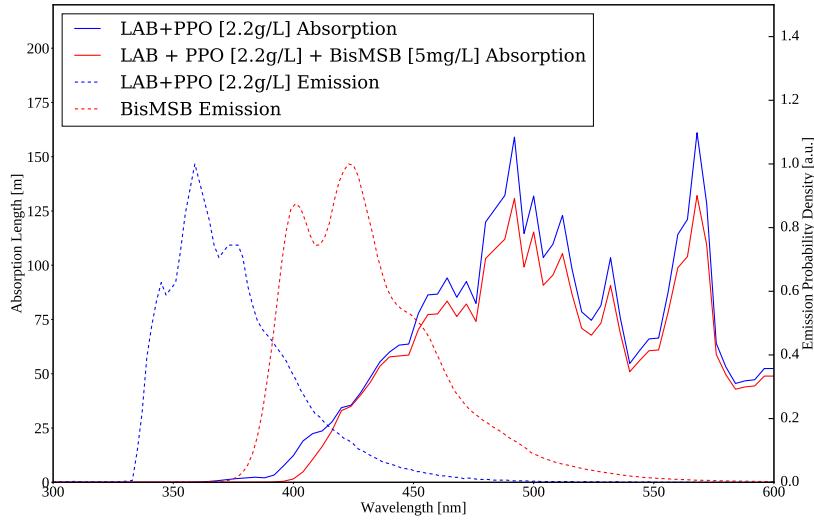
<sup>9</sup> Adding in one additional component doesn’t have to be the end, either. In SNO+  
<sup>10</sup> we plan on adding in the compound BisMSB to the scintillator cocktail. This is a  
<sup>11</sup> ‘wavelength-shifter’: scintillation light at short wavelengths is absorbed, and then  
<sup>12</sup> re-emitted at longer wavelengths, where the detection efficiency of the PMTs is greater.  
<sup>13</sup> More on the properties of the PMTs in SNO+ can be found in Section 2.3.3. The  
<sup>14</sup> net effect of the three scintillating components within SNO+ can be seen in Fig. 2.3.  
<sup>15</sup> Note how, as energy is transferred from one scintillation component to another, the  
<sup>16</sup> wavelength of light emitted gets necessarily longer as energy is lost to heat.

<sup>17</sup> The light yield of a scintillator, i.e. the amount of optical photons generated per  
<sup>18</sup> unit of energy deposited into the scintillator, is a function not just of the scintillator  
<sup>19</sup> but also the incident particle’s ionisation strength. In particular,  $\alpha$  particles are far  
<sup>20</sup> more effective at exciting and ionising nearby atoms, and so can deposit far more of its  
<sup>21</sup> energy into the scintillator per unit volume. However, the strength of this ionisation for  
<sup>22</sup>  $\alpha$ s can actually become at detriment to the generation of scintillation light. Empirically,  
<sup>23</sup> scintillators follow to first order Birks’ Law for their scintillation light yield [18]:

$$\frac{dL}{dx} = S \frac{\frac{dE}{dx}}{1 + k_{\text{Birks}} \frac{dE}{dx}}, \quad (2.3)$$

---

<sup>24</sup><sup>1</sup>To be pedantic, photons are still transferred in this energy exchange, but they are virtual instead of real.



**Fig. 2.3:** Comparison of the SNO+ liquid scintillator cocktail’s absorption and emission properties, with/without BisMSB [16, 17].

where  $\frac{dL}{dx}$  is the number of photons emitted per unit track length,  $\frac{dE}{dx}$  is the energy loss of the incident particle per unit track length,  $S$  is the scintillator’s characteristic light yield constant, and  $k_{\text{Birks}}$  is the scintillator’s “Birks’ Constant”<sup>2</sup>. For minimum-ionising particles such as a 6 MeV electron, the denominator of this equation is close to 1, and so the amount of scintillation light generated is just  $\frac{dL}{dx} = S \cdot \frac{dE}{dx}$ . However, for  $\alpha$ -particles generated in radioactive decays, this denominator can become substantial, and in the limiting case we have merely  $\frac{dL}{dx} = \frac{S}{k_{\text{Birks}}}$ . In the current phase of SNO+,  $S$  and  $k_{\text{Birks}}$  are measured to be 14,000  $\gamma/\text{MeV}$  and 0.077 mm MeV<sup>-1</sup>, respectively [19].

### 2.3.2 Optical Processes

Once optical-wavelength photons have been created within the detector, various processes can then occur that can hinder its path towards a PMT, and therefore modify the observed signal. This subsection covers the main optical processes, with a fo-

<sup>2</sup>Birks’ Constant is often just written as  $k_B$ , but this is easily confused with the far-better known Boltzmann Constant, which is completely different!

<sup>1</sup> cus on Rayleigh scattering, as an understanding of this phenomenon is critical for  
<sup>2</sup> Chapters 3–5.

### <sup>3</sup> Rayleigh Scattering

<sup>4</sup> Optical scattering is the general process of how light is scattered by particles within a  
<sup>5</sup> medium. This is fundamentally an electrodynamical process: an electromagnetic wave  
<sup>6</sup> is incident on the set of particles within the medium, which induces these particles to  
<sup>7</sup> oscillate within the field, and therefore generating their own electromagnetic radiation  
<sup>8</sup> in response. Usually, this ‘scattered’ radiation has the same frequency as that of the  
<sup>9</sup> incident radiation, and therefore the scattering is said to be *elastic*. It is possible under  
<sup>10</sup> certain circumstances for this scattered radiation to be of a longer wavelength than  
<sup>11</sup> the incident radiation: in which case, energy was absorbed by the particles and so the  
<sup>12</sup> scattering was *inelastic*. However, this latter type of scattering, also known as Raman  
<sup>13</sup> scattering, is not relevant for SNO+.

<sup>14</sup> The general solution to elastic optical scattering was first described by Gustav  
<sup>15</sup> Mie [20] and Ludvig Lorenz [21] in what is now known as *Mie Theory*. In this theory, it  
<sup>16</sup> is assumed that a plane wave of wavelength  $\lambda$  is incident on a dielectric sphere of radius  
<sup>17</sup>  $a$ . While the general solution to the problem of Mie scattering is somewhat complicated  
<sup>18</sup> (if tractable), in certain regimes one can make further simplifying assumptions that  
<sup>19</sup> substantially reduce the complexity of the result. In particular, if one assumes that  
<sup>20</sup> the size of the particle is much smaller than the wavelength of light, and that any  
<sup>21</sup> induced dipole moment can actually be established in the time window allowed by  
<sup>22</sup> the oscillation period of the electromagnetic field [22], then one can obtain *Rayleigh*  
<sup>23</sup> *scattering*. This simpler case is so-called because of its initial formulation by Lord  
<sup>24</sup> Rayleigh [23].

## 2.3 Detecting and Recording an Event in SNO+: A Journey

19

One can show that the differential cross-section associated with Rayleigh scattering of unpolarised light off of a single particle,  $\frac{d^2\sigma_{\text{Ray}}}{d\theta d\phi}(\theta, \phi)$ , is given by [24]:

$$\frac{d^2\sigma_{\text{Ray}}}{d\theta d\phi}(\theta, \phi) = \frac{8\pi a^6}{\lambda^4} \left( \frac{n_{\text{par}}^2 - 1}{n_{\text{par}}^2 + 2} \right)^2 (1 + \cos^2 \theta). \quad (2.4)$$

Here,  $\theta$  and  $\phi$  correspond respectively to the polar and azimuthal angles of the scattered waves, and  $n_{\text{par}}$  is the refractive index of the scattering particle. Most important to notice about this equation is that the cross-section follows a strong  $1/\lambda^4$  dependence, meaning that short wavelengths of light will be scattered to far greater extents than that of longer wavelengths. Secondly, the light is not scattered isotropically, but according to a  $1 + \cos^2 \theta$  dependence. This means that most light is either scattered directly forwards or backwards (known as *backscattering*), and little gets scattered orthogonally to the direction of the incident light. This is useful when it comes to trying to measure scattering in the SNO+ detector, as it provides a handle upon which to distinguish scattered light from isotropically-emitted scintillation light.

Of course, we care about the scattering that occurs within an entire bulk medium, not just the scattering off of a single molecule. From a macroscopic perspective, the key quantity of interest is a material's *Rayleigh scattering length*,  $l_{\text{Ray}}$ : the mean distance a photon is expected to travel before Rayleigh scattering. One can show that, assuming the above differential scattering cross-section, the Rayleigh scattering length is given by [25]:

$$l_{\text{Ray}} = \left[ \frac{16\pi}{3} R \right]^{-1}. \quad (2.5)$$

$R$  is the *Rayleigh ratio*,  $R = \frac{1}{V} \frac{d^2\sigma_{\text{Ray}}(90^\circ)}{d\theta d\phi}$ , equivalent to the power of the scattered light per unit volume of the scattering medium per unit incident intensity at  $\theta = 90^\circ$ .

This can lead to a few changes to Rayleigh scattering that are worth noting. Firstly, unlike for a single particle, the electric polarisability of a material can be

<sup>1</sup> *anisotropic.* Anisotropic materials have a modified angular dependence on their  
<sup>2</sup> differential cross-section, governed by the *depolarisation ratio*,  $\delta$ . In particular, the  
<sup>3</sup>  $(1 + \cos^2 \theta)$  dependence becomes  $(1 + \frac{1-\delta}{1+\delta} \cos^2 \theta)$ . For isotropic materials,  $\delta = 0$ , and  
<sup>4</sup> so the angular dependence reduces to the original form.

<sup>5</sup> Secondly, the above model has been shown to be insufficient to describe liquids  
<sup>6</sup> or solids [26], because of the non-negligible strength of their inter-molecular forces.  
<sup>7</sup> Fortunately, Einstein [27], Smoluchowski [28], and Cabannes [29] developed a theory  
<sup>8</sup> for describing how photons can scatter off of the local charge density fluctuations that  
<sup>9</sup> naturally are present in a medium because of the thermal motion of molecules. The  
<sup>10</sup> theory shows that the Rayleigh ratio of a medium is related to the medium's dielectric  
<sup>11</sup> constant,  $\varepsilon$ , by:

$$R = \frac{\pi^2}{2\lambda^4} \left[ \rho \left( \frac{\partial \varepsilon}{\partial \rho} \right)_T \right]^2 k_B T \kappa_T \frac{6 + 6\delta}{6 - 7\delta}, \quad (2.6)$$

<sup>12</sup> where  $\rho$  is the density of the medium,  $\left( \frac{\partial \varepsilon}{\partial \rho} \right)_T$  is the partial derivative of the dielectric  
<sup>13</sup> constant with respect to a changing density assuming a constant temperature  $T$ ,  $k_B$   
<sup>14</sup> is the Boltzmann Constant, and  $\kappa_T$  is the medium's isothermal compressibility. This  
<sup>15</sup> latter quantity is given by the rate of change of volume given a changing pressure of  
<sup>16</sup> the medium, all at a constant temperature.  
<sup>17</sup>

<sup>18</sup> Furthermore, the Eykman Equation [25, 30] has been shown to be an effective  
<sup>19</sup> empirical formula relating how  $\varepsilon$  is impacted by density fluctuations to the medium's  
<sup>20</sup> refractive index,  $n_{\text{med}}$ :

$$\rho \left( \frac{\partial \varepsilon}{\partial \rho} \right)_T = \frac{(n_{\text{med}}^2 - 1)(2n_{\text{med}}^2 + 0.8n_{\text{med}})}{n_{\text{med}}^2 + 0.8n_{\text{med}} + 1}. \quad (2.7)$$

<sup>21</sup> This leads to a final formula for the Rayleigh scattering length:

$$l_{\text{Ray}} = \left[ \frac{8\pi^3}{3\lambda^4} \left( \frac{(n_{\text{med}}^2 - 1)(2n_{\text{med}}^2 + 0.8n_{\text{med}})}{n_{\text{med}}^2 + 0.8n_{\text{med}} + 1} \right)^2 k_B T \kappa_T \frac{6 + 3\delta}{6 - 7\delta} \right]^{-1}. \quad (2.8)$$

---

2.3 Detecting and Recording an Event in SNO+: A Journey**21**

Discussions of the scattering lengths currently assumed within SNO+’s optical model  
 for UPW and LABPPO can be found within the theses of Krishanu Majumdar [31]  
 and Esther Turner [32]. In particular, whilst the scattering length of the UPW in the  
 water phase was measured by Esther, major systematics in the measurement remained.  
 Measurements of the scattering lengths in scintillator are a major focus of Chapters 3–5.

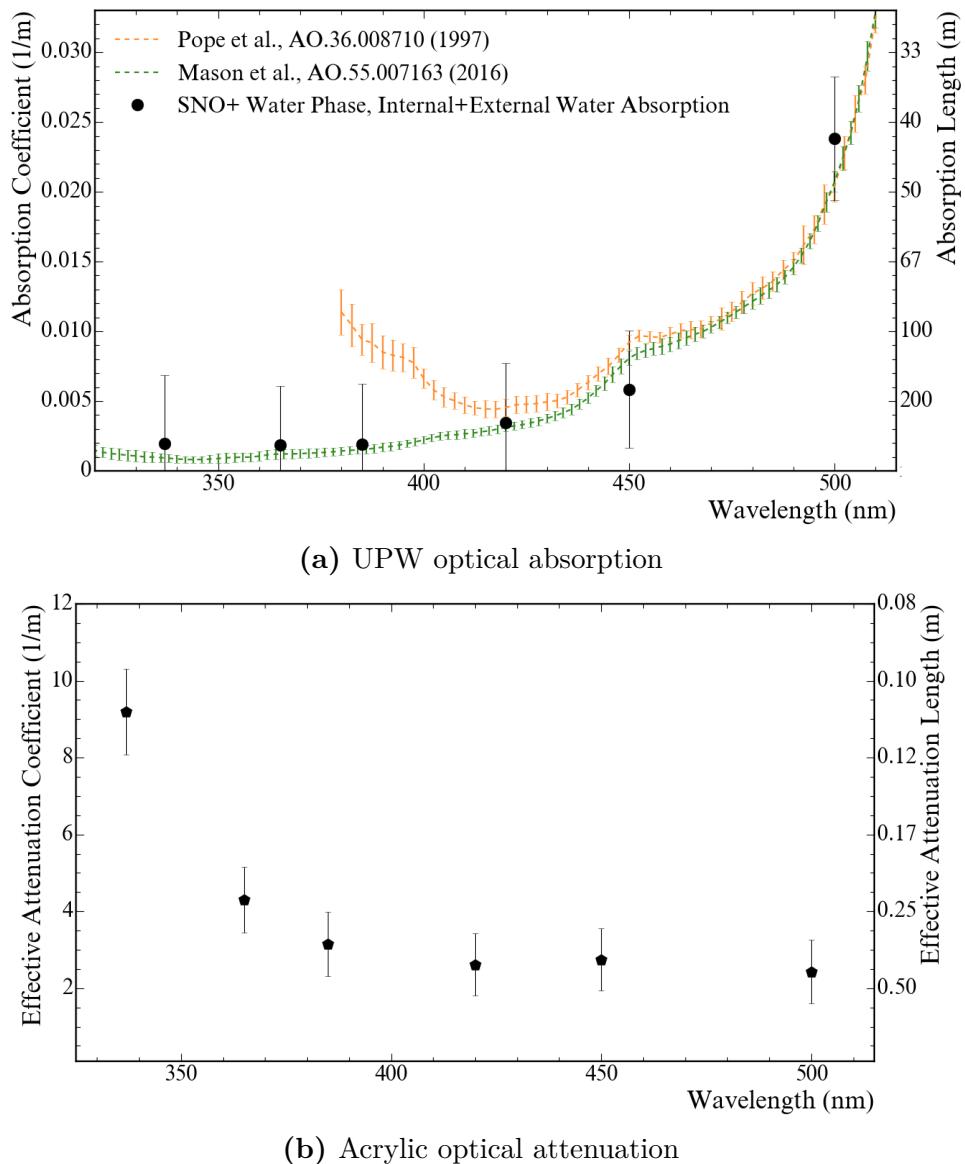
**Absorption and Re-emission**

In addition to scattering, an optical medium is also able to absorb light that propagates  
 through it. For a given medium, the *absorption length*  $l_{\text{abs}}$  is analogous to  $l_{\text{Ray}}$  described  
 above, and is typically strongly a function of wavelength. For most materials, absorbed  
 light is forever lost, converted into heat. However, for the special case of scintillators,  
 re-emission of absorbed light is possible: this is because of the physics described in  
 Section 2.3.1.

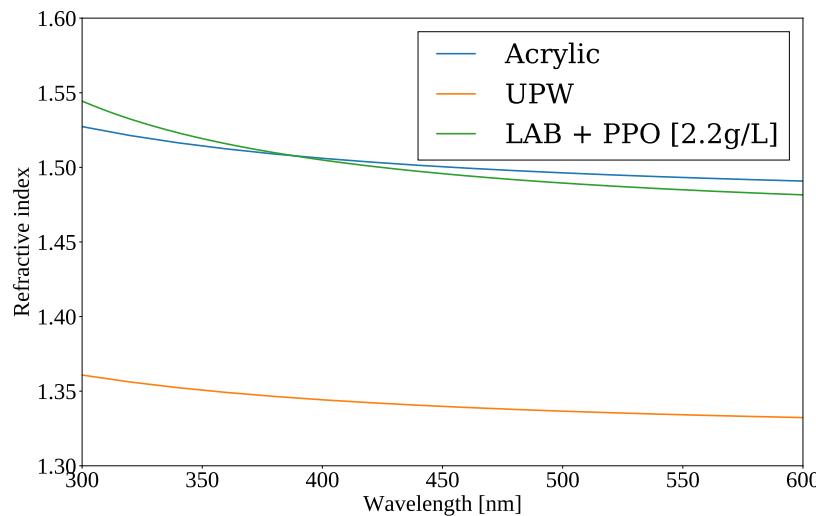
Because both scattering and absorption impede a photon’s ability to propagate  
 through a medium directly, it is often possible to measure their combined impact  
 through what is known as the attenuation/extinction length,  $l_{\text{ext}}$ :

$$\frac{1}{l_{\text{ext}}} = \frac{1}{l_{\text{abs}}} + \frac{1}{l_{\text{Ray}}}. \quad (2.9)$$

In the water phase, the ‘Laserball’ calibration system was used to measure various  
 optical properties of the detector, including the extinction lengths of the UPW and  
 acrylic as a function of wavelength [3]. Using the water phase scattering measurements  
 made by Esther, Eq. 2.9 allowed for the estimation of the absorption lengths of these  
 two materials, shown in Figure 2.4. Measurements of the extinction length in the  
 scintillator phase is discussed in detail in Chapter 5.



**Fig. 2.4:** Measured properties of the UPW and acrylic in the water phase, from [3].



**Fig. 2.5:** Refractive indices of acrylic, UPW, and LABPPO as a function of wavelength [3, 33, 34].

### Surface reflection and refraction

When light travels through the boundary of one medium to another, both reflection and refraction can be possible, depending on the relative refractive indices of the two media. The refractive indices of the UPW, acrylic, and LABPPO are shown as a function of wavelength in Figure 2.5. Note that, for most optical wavelengths, LABPPO has a very close refractive index to acrylic, whereas UPW is somewhat farther away. By consequence, negligible refraction is expected in most cases for light travelling between the liquid scintillator and the acrylic; however, substantial refraction is possible for light travelling between acrylic and UPW. Because of this, isotropically-emitting point-like physics events within the AV that are close enough to the acrylic will have some of their light undergo Total Internal Reflection (TIR) at the AV, reflecting back into the AV instead of continuing outward into the outer water.

Even when not undergoing TIR, some light at a boundary can still reflect. The fraction of light that reflects is known as the *reflectance*  $R$ , compared to that which is able to transmit through the boundary, the *transmittance*  $T = 1 - R$ . The *Fresnel*

<sup>1</sup> Equations determine the reflectance of an interface [35]:

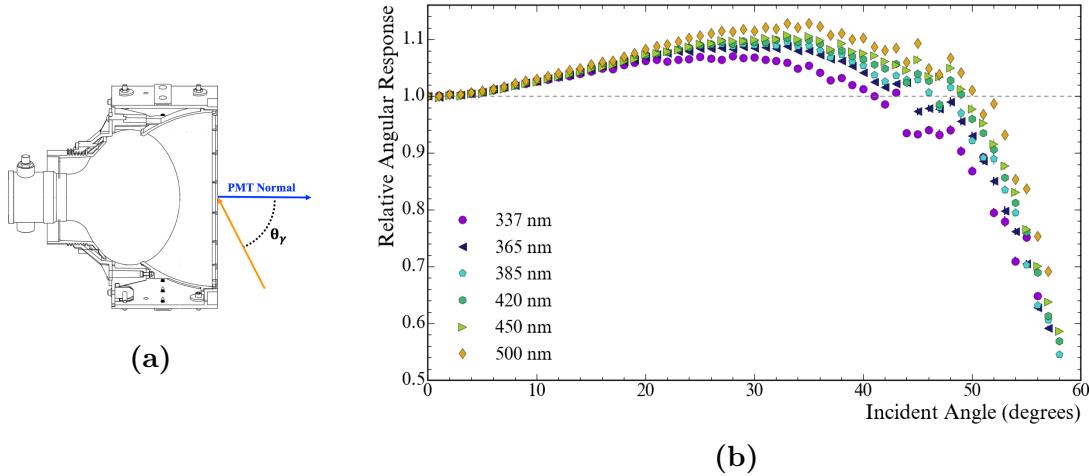
$$\text{where } R_s = \left| \frac{n_1 \cos \theta_i - n_2 \cos \theta_t}{n_1 \cos \theta_i + n_2 \cos \theta_t} \right|^2, R_p = \left| \frac{n_1 \cos \theta_t - n_2 \cos \theta_i}{n_1 \cos \theta_t + n_2 \cos \theta_i} \right|^2, \quad (2.10)$$

<sup>3</sup> where  $R_s$  and  $R_p$  are the reflectances of *s*- and *p*-polarised light,  $n_1$  and  $n_2$  are the  
<sup>4</sup> refractive indices of the first and second optical media, and  $\theta_i$  and  $\theta_t$  are the angles of  
<sup>5</sup> incidence and refraction, respectively. For SNO+, we are only interested in unpolarised  
<sup>6</sup> light, so the total reflectance  $R = (R_s + R_p) / 2$ .

### <sup>7</sup> 2.3.3 Detection by PMTs

<sup>8</sup> The final step for photons in our journey is detection by a PMT. Almost all PMTs in  
<sup>9</sup> SNO+ are of the Hamamatsu R1408 design [1]. These PMTs within SNO+ are housed  
<sup>10</sup> within a reflecting cone known as a ‘concentrator’. The combined PMT–concentrator  
<sup>11</sup> ‘bucket’, shown in Fig. 2.6a, is designed to maximise the collection efficiency of light  
<sup>12</sup> emanating from within the AV, whilst minimising the collection efficiency of light  
<sup>13</sup> outside the AV [36]. The so-called ‘angular response’ of the PMT buckets has been  
<sup>14</sup> measured in both SNO and SNO+ using the Laserball, which describes the relative  
<sup>15</sup> collection efficiency as a function of the polar angle of the incident light ray relative  
<sup>16</sup> to the direction in which the PMT bucket points. The results of this can be seen in  
<sup>17</sup> Fig. 2.6b.

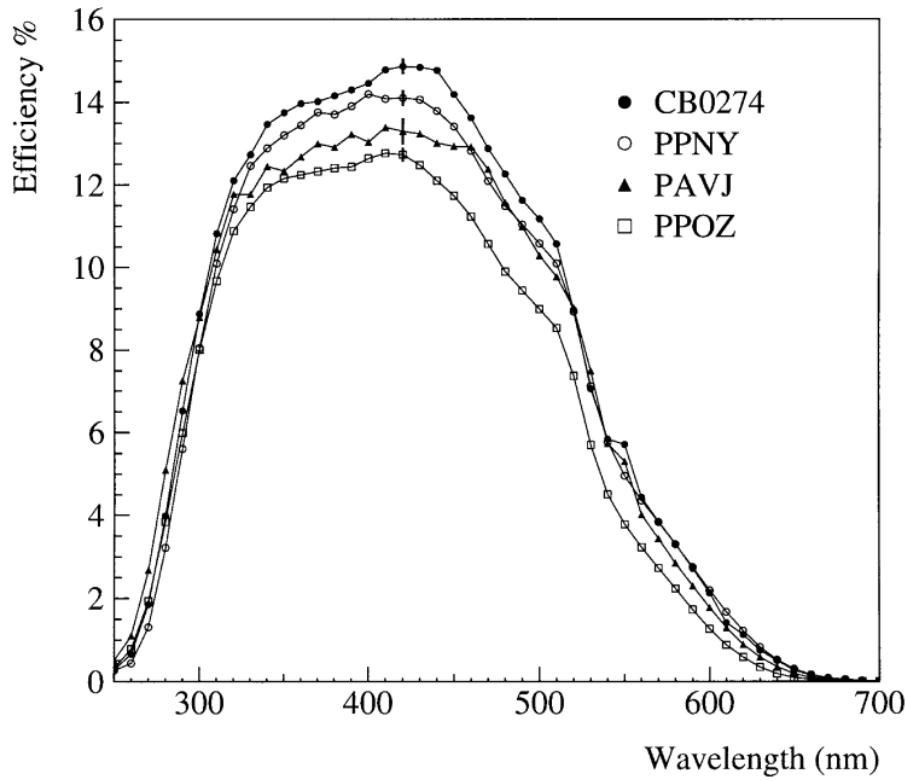
<sup>18</sup> Once a photon is incident on the PMT’s photocathode, it is possible for that photon  
<sup>19</sup> to be absorbed and generate a photoelectron. The probability of this happening is  
<sup>20</sup> governed by the photocathode’s Quantum Efficiency (QE) at the photon’s wavelength,  
<sup>21</sup> as well as the collection efficiency of a photoelectron onto the first dynode of the  
<sup>22</sup> PMT. The combined measured efficiency of PMTs tested *ex-situ* for SNO can be  
<sup>23</sup> seen in Fig 2.7. Once this photoelectron has been created, the dynodes within the  
<sup>24</sup> PMT generate a cascade of electrons that allow for an observable voltage signal to



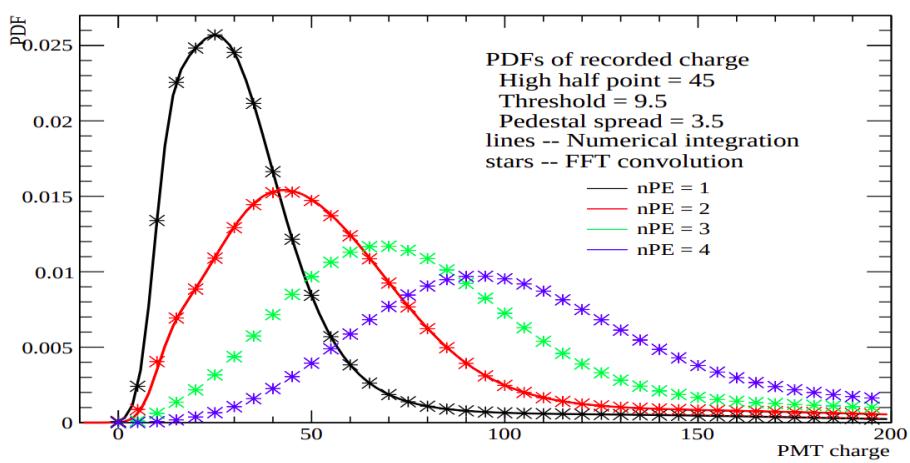
**Fig. 2.6:** (a): Diagram of the PMT and concentrator ‘bucket’ used within SNO+, showing also the definition of the incidence angle. (b): Plot of the measured relative angular response of the PMTs in SNO+, as a function of both incidence angle and wavelength. Both figures taken from [3].

be produced. The dynamics of this cascade are such that there is a natural spread of possible times between the creation of a photoelectron and the generation of the voltage signal pulse in the PMT’s wire. This is known as the ‘Transit Time Spread’ (TTS) of the PMTs: for SNO+, the RMS timing resolution of the TTS for the R1408-type PMTs is 1.7 ns [1].

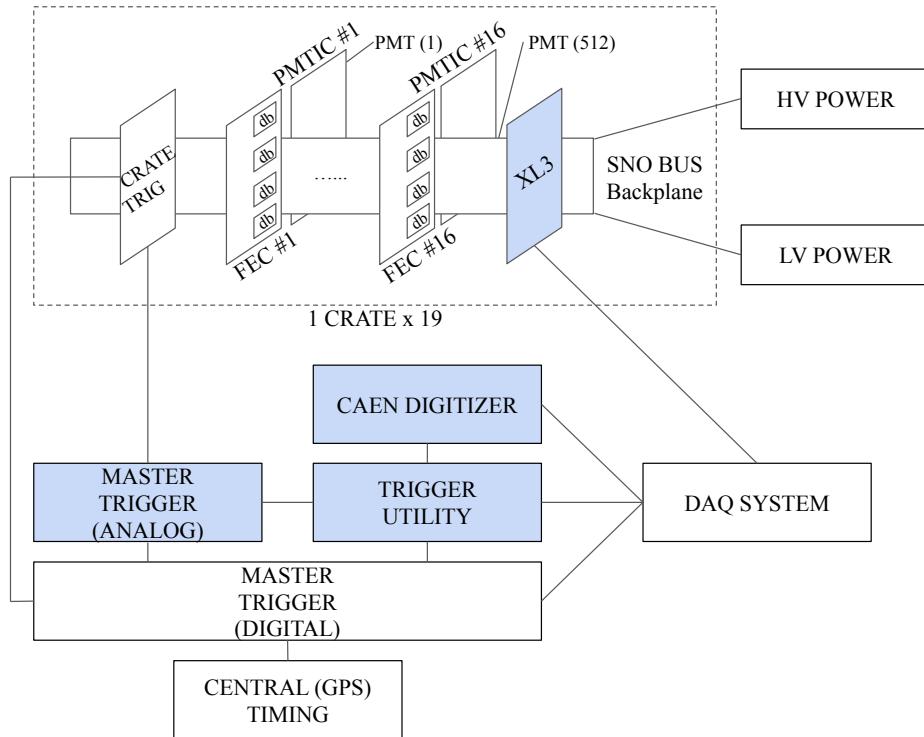
Finally, if multiple photons generate photoelectrons on the same PMT close enough in time, the detector’s front-end electronics (described in Section 2.3.4) is unable to tell. However, the amount of charge generated increases in proportion to the number of photoelectrons (npe). Much like with the transit time, the strength of the signal observed by the PMT is governed by a distribution, a function of the npe generated. Examples of these distributions can be seen in Fig. 2.8. The relatively large widths of these charge distributions precludes the ability to straightforwardly determine the npe purely from charge when the npe is small. To work around this, various techniques can be employed to try and estimate the npe in a given PMT — an example of one such method can be seen in Section 4.2.



**Fig. 2.7:** Efficiencies of the R1408-type PMTs used as standard within SNO+ [37].



**Fig. 2.8:** Example charge spectra for a PMT as a function of the true npe generated [38].



**Fig. 2.9:** Schematic of the front-end electronics used for data acquisition and triggering in SNO+, taken from and discussed in [2].

### 2.3.4 Data Acquisition and Triggering

Once a signal is generated in the cable attached to a PMT, it travels along up to the front-end electronics on the deck above the detector. The job of these electronics, known as the data acquisition (DAQ) and triggering system, is to convert raw electronic signals from the PMTs into recorded digital ‘events’ that can be used for analysis. A schematic showing the setup of the electronics is shown in Fig 2.9, with full details in [2].

A signal passes first through the PMT Interface Card (PMTIC), which then sends it through to one of the Daughter Boards (DBs) which are stored on Front-End Cards (FECs) within one of 19 electronic crates on deck. The DBs determine if the analogue signal along a given PMT channel has crossed a pre-defined charge threshold, at which

<sup>1</sup> point we say that we have detected a ‘hit’ on that PMT’s channel. When this occurs,  
<sup>2</sup> the DB performs a set of important actions:

- <sup>3</sup> 1. Begins a timer for that channel, in the form of a Time-to-Amplitude Converter  
<sup>4</sup> (TAC). TAC also corresponds to the resulting quantity being measured.
- <sup>5</sup> 2. Begins integrating the total charge signal for that channel in three ways, known  
<sup>6</sup> as QHS, QHL, and QLX. These correspond to using different integration times  
<sup>7</sup> and gain settings.
- <sup>8</sup> 3. Generates trigger pulses for that channel for each available trigger type. Three  
<sup>9</sup> main trigger signals are the ‘N20’ (a square pulse for 20 ns), ‘N100’ (a square  
<sup>10</sup> pulse for 100 ns), and ‘ESUMHI’ (a pulse copying the shape of the voltage signal  
<sup>11</sup> for that channel). The reason for these names will be explained shortly.

<sup>12</sup> Whilst the TAC, QHS, QHL, and QLX are being calculated, the trigger signals  
<sup>13</sup> from each channel are sent over to the Crate’s Trigger Card (CTC), where the signals  
<sup>14</sup> are then summed for each trigger type. These crate-level trigger signals are then sent  
<sup>15</sup> over to the 7 detector-level Analogue Master Trigger Cards (MTC/A+), which further  
<sup>16</sup> sum the signals by trigger type from all the crates in the detector. It is at this point of  
<sup>17</sup> the process where the names of the trigger types becomes clear: the combined N20  
<sup>18</sup> and N100 signals are proportional to the total number of hit PMTs within a 20 ns and  
<sup>19</sup> 100 ns time window, respectively, whilst the total ESUMHI signal corresponds to the  
<sup>20</sup> total charge seen over all the PMTs.

<sup>21</sup> If these trigger signals go above certain pre-defined thresholds, then a signal is  
<sup>22</sup> sent for that trigger type to the Digital Master Trigger Card (MTC/D). The MTC/D  
<sup>23</sup> receives all trigger signals from the detector, and if a given trigger type has been  
<sup>24</sup> ‘masked in’ the Card will generate a Global Trigger (GT) for the detector according to  
<sup>25</sup> its 50 MHz clock. Under certain circumstances, such as calibrations, a trigger signal can

---

2.3 Detecting and Recording an Event in SNO+: A Journey

29

be generated externally and asynchronous to the MTC/D: these are ‘EXTA’ triggers.  
1  
Any such EXTA trigger signal is first handled by an electronics box named ‘TUBii’  
2  
(Trigger Utility Board Mark ii), before then being passed onto the MTC/D. This  
3  
becomes relevant when discussing the calibration electronics described in Chapter 3.  
4

Once a GT signal is generated, it is then sent back to all the CTCs, which then  
5  
orders the integration of time and charge to be stopped on all channels for that crate.  
6  
The time and charge information that has been temporarily stored on each channel’s  
7  
CMOS chip on the FEC is then sent to the crate’s ‘XL3’ card, which packages the  
8  
crate’s raw information via ethernet over to a set of computers. The trigger signals for  
9  
a triggered event are also digitised by a CAEN brand Analogue-to-Digital Converter  
10  
(ADC), and sent to the same DAQ computers. The total window of time in which  
11  
data is gathered from one GT signal is 400 ns, with data from up to 180 ns before and  
12  
220 ns after the GT has arrived. There is then a necessary ‘dead time’ of 420 ns after a  
13  
given GT has been given in which no further GTs can be made.  
14

Finally, the raw data from the crates and trigger system arrives in a set of computers,  
15  
which organise all of this into an individually-packaged ‘event’, stored on disk in the  
16  
'Zebra Database' (ZDAB) format. A given built event contains the TAC and charge  
17  
information from each hit PMT, the CAEN digitised waveforms, a unique identifying  
18  
number for that triggered event (the GTID), as well as the times from both the  
19  
MTC/D’s 50 MHz clock and a GPS-calibrated 10 MHz clock. The former time is used  
20  
for measuring relative times between events whilst the latter is used for knowing the  
21  
time of day of an event: both are used in Chapter 6.  
22

### 2.3.5 Operation of the Detector

Control of the detector’s DAQ system is handled through a custom-built GUI known  
24  
as ORCA [39]. This program allows operators of the detector to modify settings in the  
25

<sup>1</sup> detector electronics at both a low- and high-level. It also allows operators to monitor  
<sup>2</sup> the current status of the detector, such as voltage and current levels within each crate.

<sup>3</sup> ORCA allows for the detector to have its data split into ‘runs’ of different types.

<sup>4</sup> For the majority of the time, the detector is run in the ‘Physics’ mode, with individual  
<sup>5</sup> runs split into 1 hour periods. It is this data that is used for almost all high-level  
<sup>6</sup> physics analyses, such as the one described in Chapter 6. Other modes include ones  
<sup>7</sup> for detector maintenance, as well as for calibrations of various kinds. During certain  
<sup>8</sup> calibration runs, data can be further split into ‘subruns’ where necessary. Operation of  
<sup>9</sup> specific calibration sources, including the one described in Chapter 3, can be performed  
<sup>10</sup> through the ORCA GUI.

## <sup>11</sup> **2.4 Calibrations and Detector Modelling**

<sup>12</sup> Once the raw data from triggered events has been stored in files, certain extra steps  
<sup>13</sup> must be taken before effective analysis of that data can be achieved. This section  
<sup>14</sup> covers those steps.

### <sup>15</sup> **2.4.1 Detector Monitoring**

<sup>16</sup> No data taken from the detector can reasonably be used for analysis unless its quality  
<sup>17</sup> has been approved. This is done in a number of ways on SNO+. Firstly, a number of  
<sup>18</sup> automated systems monitor all aspects of the detector, including voltage levels in the  
<sup>19</sup> crates, trigger rates, as well as ‘slower’ quantities such as the tensions on the ropes  
<sup>20</sup> holding the AV in place. Problems in any of these measured parameters trigger an  
<sup>21</sup> automatic alarm system, which notifies a human detector operator. A human detector  
<sup>22</sup> operator monitors the detector 24/7 whilst the detector is live.

In addition to systems that monitor whether anything has gone wrong, information about the state of the detector during each run is stored in a database known as RATDB. This information includes, amongst other things, a recording of which PMT channels have actually been raised to high voltage for that run, as well as any channels/cards/crates that have been flagged for having a known poor data quality (e.g. being overly noisy).

### 2.4.2 Electronic and PMT Calibrations

The lowest level of calibrations performed in SNO+ are the Electronic and PMT Calibrations: ECAs and PCAs, respectively. These calibrations convert the raw time and charge values recorded by the DAQ into quantities that can actually be used in analysis.

During an ECA, two main quantities are measured. Firstly, because of noise the integrated charge measured on each channel is offset by some amount. This offset, known as the ‘pedestal’, is recorded for each channel. The other quantity is the ‘time slope’ for each channel, which allows one to convert from the ADC TAC counts into an uncalibrated hit time of that channel’s PMT in ns. Both of these quantities are measured by manually sending external signals to channels in the crates, forcing them to start measuring TAC and charge even though no PMTs were actually hit. Running ECAs also enable us to spot any channels with unusual behaviour, so that they are not used during analysis.

Using ECAs alone is not enough to have fully-calibrated time and charge data. The lengths of cables between PMTs and PMTICs are all slightly different, leading to differences in the so-called ‘cable delay’ of each channel. This means that two PMTs that have a photoelectron generated at the same time can generate slightly different TAC values. Furthermore, because the start time of the TAC is determined by when

<sup>1</sup> the channel’s signal goes above a constant threshold, if a signal is very large (e.g. when  
<sup>2</sup> numerous photoelectrons have been generated on one PMT) then the start time of the  
<sup>3</sup> TAC will be systematically earlier. This is known as the ‘time walk’. Both of these  
<sup>4</sup> quantities get measured during PCAs.

<sup>5</sup> PCAs can be performed by either the Laserball or by the TELLIE calibration  
<sup>6</sup> system. The latter is a series of 92 optical fibres attached at various points of the  
<sup>7</sup> PSUP, through which optical-wavelength light can be fired from LEDs. TELLIE is the  
<sup>8</sup> Timing subsystem of the ELLIE calibration system: the Embedded LED/Laser Light  
<sup>9</sup> Injection Entity. The other two subsystems, AMELLIE and SMELLIE, are introduced  
<sup>10</sup> in Section 2.4.3. For both the Laserball and TELLIE calibration systems, the cable  
<sup>11</sup> delay and time walk are measured by firing light from the source at a known time, and  
<sup>12</sup> observing when the signal arrives in each PMT channel.

<sup>13</sup> On top of calibrating the PMT hit times, PCAs also further calibrate the charge  
<sup>14</sup> information. In particular, the charge spectrum generated by a single photoelectron is  
<sup>15</sup> determined for each channel. This allows us to convert the pedestal-corrected charge  
<sup>16</sup> ADC counts into a number of photoelectrons.

<sup>17</sup> Using the data gathered from both ECAs and PCAs, the raw data stored in  
<sup>18</sup> ZDABs is processed into a new file format known as RATDS files. These files contain all  
<sup>19</sup> the information of an event, but now the timing and charge information have been  
<sup>20</sup> calibrated. It is this file type used in the optical calibration work of Chapters 3–5.

### <sup>21</sup> 2.4.3 Energy and Optical Calibrations

<sup>22</sup> The next stage of calibrating the detector is modelling its optical properties. These  
<sup>23</sup> properties include all the processes covered in Section 2.3.2, such as scintillator emission,  
<sup>24</sup> optical absorption, re-emission, and Rayleigh scattering. This is crucial, as it allows us to

---

reconstruct information about events within the detector: more on event reconstruction  
shortly.

In addition to deployments of the Laserball (discussed in Section 2.3.2), two further  
calibration sources are used in SNO+ to measure properties of light propagation:  
AMELLIE and SMELLIE. These are the ‘Attenuation Module’ and ‘Scattering Module’  
for the ELLIE calibration system. Like TELLIE, AMELLIE and SMELLIE consist of  
optical light sources that shine through optical fibres into the detector. The former  
uses LEDs from TELLIE, whilst the latter uses optical wavelength lasers. Despite the  
names both subsystems are similar enough that they are both capable of measuring  
attenuation and scattering within the detector. More details about the SMELLIE  
hardware can be read in Chapter 3.

Another critical component of the detector to calibrate well is the energy response:  
given a specific amount of energy deposited in the water/scintillator, how many hits are  
observed? For this, a number of radioactive sources are used at a variety of energies. In  
the scintillator phase, there are three main sources. The first is an americium-beryllium  
(AmBe) source inherited from SNO [40], which contains  $^{241}\text{Am}$  that  $\alpha$ -decays, which  
can be captured by the  $^9\text{Be}$  within the source. This capture leads to the emission of a  
neutron as well as production of a  $^{12}\text{C}$  nucleus, which 60% of the time is in an excited  
state. When this excited state decays, a 4.4 MeV  $\gamma$  is emitted promptly. Eventually the  
neutron is captured by hydrogen in the detector, leading to a characteristic 2.2 MeV  
delayed  $\gamma$  being generated [2]. Both the prompt and delayed peak energies from the  
AmBe source can be used for energy calibration. In addition to this, one can calibrate  
the neutron detection efficiency with the AmBe source [40], which is important for the  
analysis of antineutrino IBD events.

Another deployable radioactive source is the  $^{16}\text{N}$  source, also originally used for  
SNO [41]. The  $^{16}\text{N}$  isotope  $\beta$ -decays to  $^{16}\text{O}$ , with a distinctive 6.1 MeV  $\gamma$  also being

<sup>1</sup> generated 66% of the time. It is the  $\gamma$  that can make it out to the detector, whilst the  
<sup>2</sup>  $\beta$  can be tagged by a block of scintillator and PMT held within the source container.

<sup>3</sup> Neither of the above two calibration sources have been deployed internally within  
<sup>4</sup> the scintillator during the scintillator phase because of concerns over contaminating  
<sup>5</sup> the scintillator. However, there have been a number of deployments in the external  
<sup>6</sup> water. Alongside this, a different kind of radioactive source has been used during the  
<sup>7</sup> scintillator phase for energy calibration: the existing radioactive background spectra  
<sup>8</sup> within the detector. Backgrounds such as  $^{14}\text{C}$ ,  $^{210}\text{Po}$ ,  $^{214}\text{BiPo}$ , and  $^{208}\text{Tl}$  all have  
<sup>9</sup> distinctive peaks in the energy spectrum of SNO+ during the scintillator phase, and  
<sup>10</sup> can be used to calibrate the scintillator’s energy response. Using  $^{214}\text{BiPo}$  events in  
<sup>11</sup> particular has been used for energy scale calibration with the solar oscillation analysis,  
<sup>12</sup> as discussed in Section 6.2.3.

#### <sup>13</sup> 2.4.4 Event Reconstruction

<sup>14</sup> Once the detector has been calibrated, event ‘reconstruction’ becomes possible. This is  
<sup>15</sup> the process of deriving high-level physics quantities about a triggered event within the  
<sup>16</sup> detector, based upon the calibrated hit information. In SNO+, our base assumption in  
<sup>17</sup> most cases of event reconstruction is that a triggered event was due to a single point-like  
<sup>18</sup> electron track. Reconstructing an event involves running a number of algorithms, which  
<sup>19</sup> in the scintillator phase are together called the **ScintFitter**.

<sup>20</sup> The first critical pieces of information that gets determined by **ScintFitter** is the  
<sup>21</sup> event’s position and time. The reconstructed position corresponds to the point in the  
<sup>22</sup> detector where the triggered event most likely came from (assuming the event was  
<sup>23</sup> point-like in extent), whilst the reconstructed time is the starting emission time of the  
<sup>24</sup> event, relative to the event’s trigger time. The position of an event is critical to know,  
<sup>25</sup> as far fewer background events occur near the centre of the detector compared to the

edges. It is also important to know the emission time of an event, as this allows us to build the so-called ‘time residual’ ( $t_{res}$ ) distribution of an event. For a point-like physics event in the detector,  $t_{res}$  for a given PMT hit is defined as:

$$t_{res} = t_{\text{hit}} - t_{\text{TOF}} - t_{\text{emm}}, \quad (2.11)$$

where  $t_{\text{hit}}$  is the calibrated hit time of the PMT,  $t_{\text{TOF}}$  is the time one expects for light to travel directly from the reconstructed position to that PMT (the time-of-flight), and  $t_{\text{emm}}$  is the reconstructed emission time.

Whilst a number of algorithms have been developed for reconstructing position and time in SNO+, they all work on the same basic principle. Because of the spherical symmetry of the detector, if an event occurs at the centre of the detector one expects direct light to hit PMTs throughout the detector at the same time<sup>3</sup>. However, if an event happens some distance away from the detector’s centre then direct light will arrive at the PMTs it is closer to sooner. Therefore, by looking at the distribution of hit times for PMTs that were hit earliest as a function of the PMTs’ positions in the detector (ignoring PMT hits that arrived much later, presumably because the photon paths were not direct) one can try and estimate where the position of the event was. Reconstructed positions and reconstructed times are deeply linked by the time residual equation described above.

Currently on SNO+, a likelihood-based approach is used to reconstruct position and time. The algorithm endeavours to maximise the combined likelihood of the observed calibrated hit times of the hit PMTs, given proposed points in the four-dimensional (position, time) parameter space [42, 43, 44]. However, regardless of algorithm there are two factors that limit the position and timing reconstruction of an event. The

<sup>3</sup>It is possible for the centre of the AV and PSUP to not be completely aligned, i.e. there is an ‘AV offset’. Then, the spherical symmetry can be broken by refraction through the AV. Fortunately, we account for this when coordinating our position fitters.

<sup>1</sup> TTS of the PMTs used as well as the speed of the scintillator emission timing defines  
<sup>2</sup> the fundamental timescale — and hence also length scale — by which events can  
<sup>3</sup> be reconstructed. Secondly, if more photons are able to generate prompt hits in the  
<sup>4</sup> detector from a given event, then more information can be used to determine the  
<sup>5</sup> position and time. Under current conditions, a 2.5 MeV event in the centre of the  
<sup>6</sup> detector will have a position resolution of 100 mm [45].

<sup>7</sup> The other critical piece of reconstructed event information is the event’s energy.  
<sup>8</sup> More precisely, this is the kinetic energy of an electron that has been assumed to have  
<sup>9</sup> generated the event. By consequence, events due to  $\alpha$ -decay will obtain reconstructed  
<sup>10</sup> energies well below the actual energy of the  $\alpha$  particle, because of scintillator quenching.

<sup>11</sup> At its simplest, assuming that an event is from an electron of moderate energy,  
<sup>12</sup> then we can expect the number of PMT hits observed to be directly proportional to  
<sup>13</sup> the energy of the event. Given that the number of hits observed in an event (called the  
<sup>14</sup> `nhit`) is governed by a Poisson distribution, then the uncertainty in energy will just be  
<sup>15</sup> proportional to the square root of the number of hits. As a result, the reconstructed  
<sup>16</sup> energy resolution in SNO+ is determined by the scintillator’s light yield and absorption  
<sup>17</sup> length, as well as the coverage and QE of the PMTs.

<sup>18</sup> There are second-order corrections to the energy reconstruction that need to be  
<sup>19</sup> considered to minimise bias. At low energies, scintillator quenching becomes non-  
<sup>20</sup> negligible, so an understanding of the scintillator’s Birks’ constant is needed. At  
<sup>21</sup> high energies, many PMTs will have had multiple photoelectrons generated, so merely  
<sup>22</sup> using the `nhit` will give an underestimate of the true energy. Finally, the detection  
<sup>23</sup> efficiency of photons is non-uniform as a function of position in the detector. The  
<sup>24</sup> current energy reconstruction algorithm used within SNO+ attempts to deal with all  
<sup>25</sup> of these effects [46, 47].

After position, time, and energy reconstruction, **ScintFitter** calculates a number of additional quantities from what are known as classifiers. These describe a wide number of properties about an event, often using the derived  $t_{res}$  distribution of an event as the basis for classification. Some examples of classifiers used in analysis are discussed in Section 6.1.2.

All Physics data runs, as well as certain calibration data such as AmBe and  $^{16}\text{N}$ , has the **ScintFitter** algorithm run over it after having been processed for time and charge calibration. This results in what is known as a fully-processed RATDS file, as well as a new file type known as an **ntuple**. This latter file type has much of the hit-level information removed, and contains only event-level information such as the reconstructed energy and position. Because these files are much smaller, they are the ones typically used in the high-level physics analyses on the experiment, such as the one described in Chapter 6.

#### 2.4.5 Event Simulation

Simulations of events in SNO+ are performed using the software **RAT** [2]. Built on the **GEANT4** particle physics software framework, **RAT** is capable of simulating all aspects of the physics of an event within the detector via a Monte Carlo (MC) approach. This includes any particle physics that defines an event’s generation, propagation and interactions of those particles in the detector media, the generation of light by both scintillation and Cherenkov processes, the propagation of that light, as well as the detection of that light by PMTs and simulation of the expected DAQ response. **RAT** is then used to process both simulation and data in the same way. In addition to being highly customisable, **RAT** can use the **RATDB** tables generated from a given data run when simulating to try and match those particular run conditions as closely as possible. **RAT** also offers a suite of tools to assist with analysis of data. The software is

- <sup>1</sup> constantly being updated with new features — the work done in this thesis uses RAT
- <sup>2</sup> versions between 6.18.8 and 7.0.9, inclusive.

## Chapter 3

1

# The SMELLIE Calibration System

2

*There's a certain Slant of light,*

*Winter Afternoons —*

*That oppresses, like the Heft*

3

*Of Cathedral Tunes —*

---

EMILY DICKINSON

- Basic principle for how SMELLIE works: firing collimated laser light into detector to observe scattering events.
- Analysis will measure and monitor scattering in a detector with changing optics.
- One can try and measure some component of this: the cross-section/scattering length versus wavelength and time, and/or the relative scattering length versus wavelength and time.

4

5

6

7

8

9

[1 page]

10

### 3.1 The SMELLIE Hardware

- Describe the existing hardware, post-upgrade made in Summer 2022. For pre-upgrade hardware, can simply cite previous SMELLIE theses. This includes the path of light into the detector, as well as the path of the trigger signal.
- Make sure to mention explicitly these upgrades: Tony Zummo's fix to the TUBii trigger logic, as well as the addition of the VFA, updated MPU, and modified trigger window. Make sure to motivate why these updates were made.

[7 pages]

### 3.2 Software for SMELLIE Data-taking

- Can be brief here! Little has changed since previous theses, so can mostly just summarise and cite.
- Server running on SNODROP machine, which converts high-level commands into low-level ones that the hardware can interpret.
- Run plan files written in JSON handed to ORCA which then sends relevant commands to SNODROP which fires as appropriate.
- Operator interacts with ORCA to perform SMELLIE calibration runs.
- After SMELLIE data taken, run description file created, containing metadata about the run conditions, used in analysis.

[2 pages]

### 3.3 Commissioning SMELLIE in the

#### Scintillator Phase

- Explain why commissioning of SMELLIE is needed: Need to confirm that SMELLIE is working as expected; determine intensity "set-points" for different use cases.
- Commissioning originally performed by Esther and JeffL back in the water phase; explain why this needed to be re-done for both the scintillator phase and after the hardware upgrades.
- No need to describe the Tesseract in detail here - that can be in Jeff L's thesis. But, I do want to show the results of both commissioning campaigns in scintillator-fill, one before the new hardware was added, and one after.

[5 pages] [15 PAGES TOTAL]



# Chapter 4

1

## Simulating SMELLIE Events

2

Max Power : *Kids. From now on there are three ways of doing things: the right way, the wrong way, and the Max Power way.*

Bart Simpson : *Isn't that just the wrong way?*

Max Power : *Yes, but faster!*

---

THE SIMPSONS

Critical to extraction of scattering information from SMELLIE data is an accurate Monte Carlo (MC) simulation of the SMELLIE system. By modelling the laser light emission into the detector correctly, we can simulate how SMELLIE light will be impacted by changing scattering lengths in the detector. Because of the complexity of the optics of the optical fibres used to direct the laser light into the detector, a given SMELLIE event is simulated as a partially-collimated “flash” of visible photons emanating from the emission point of the fibre into the detector. This flash then requires a number of parameters to be correctly described. In particular:

4

5

6

7

8

9

10

11

- **Fibre emission positions** were recorded during the installation of the fibres.

12

- **Wavelength and emission timing distributions** of light pulses were taken from measurements of the laser heads by their manufacturers [], or by colleague Jeff Lidgard in the case of the SuperK wavelength distribution [].
- **The “pulse magnitude”**, defined as the mean number of photons simulated per event, is determined on a subrun-by-subrun basis, and is assumed to fluctuate as a Poisson distribution.
- **The beam profiles**, which describe the angular emission distributions of each fibre, is the focus of this chapter. These are necessary because unlike scintillation light, the light emitted from SMELLIE fibres is not isotropic.
- **Nominal fibre emission directions** attempt to define the centre of the beam for a given fibre.

This chapter is split into three sections. Improvements to the existing simulation algorithm for the beam profiles are first made, and then the beam profiles themselves are updated. Finally, comparisons between data and simulation are made after the upgrades to investigate any remaining discrepancies.

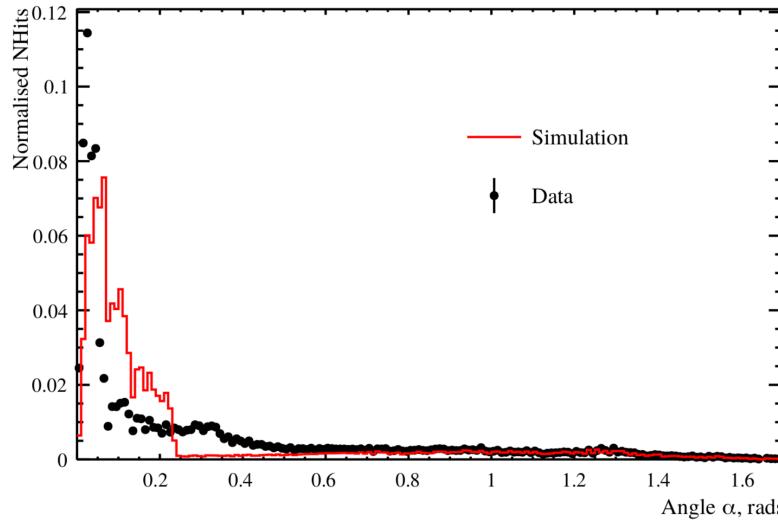
## 4.1 Improving the SMELLIE Generator Algorithm

### 4.1.1 Previous Attempts at SMELLIE Event Simulation

Before we can determine the beam profiles, we must first decide how to specify them. Previous observations show that different fibres can have notably different beam profiles [31], so we let each fibre’s beam profiles be unique. We assume for now that a given fibre’s beam profile is stable over time, and independent of the wavelength of light fired. A straightforward, naïve approach to parameterising a beam profile would be as follows: specify some nominal fibre direction, corresponding to the direction light

## 4.1 Improving the SMELLIE Generator Algorithm

## 45



**Fig. 4.1:** Comparison between a simulation of one of the fibres, made from the 1D beam profile generator (red), with the associated data subrun that was used to create that beam profile (in black). For both MC and data, what is plotted is the PDF of observed PMT hits, as a function of the  $\alpha$  angle. Poissonian errors have been added to the data points, but are too small to see. Clearly, this 1D generator does not replicate the observed beam profile correctly. Figure taken from [32].

takes travelling from the fibre to the centre of the “beamspot” observed on the other side of the detector. Then, specify a 1D beam profile, corresponding to the probability density of firing a photon at a given polar angle  $\alpha$  relative to the nominal direction. One might even assume this distribution is Gaussian. The distribution in azimuthal direction,  $\phi$ , is assumed to be uniform.

This 1D beam profile approach was used initially for SMELLIE, and remains in use for the other ELLIE sub-systems within SNO+. However, when SMELLIE data was taken in the water-phase of the experiment, simulations using these beam profiles failed to match them well at all - see figure 4.1 for an example. Not only was the distribution in  $\alpha$  not Gaussian, a distinct speckle-pattern can be observed within the beamspot that is not uniform in  $\phi$ . This fact led to colleague Esther Turner building a SMELLIE generator that could handle 2D beam profiles: dependent on both  $\alpha$  and  $\phi$ . The distribution was stored as a map from each inward-pointing PMT in the detector

<sup>1</sup> to a relative intensity value. This was chosen because the beam profile shapes were  
<sup>2</sup> calibrated from existing SMELLIE data — more on this in section 4.2.

<sup>3</sup> This original 2D generator then sampled the beam profile via a rejection sampling  
<sup>4</sup> approach, outlined as follows:

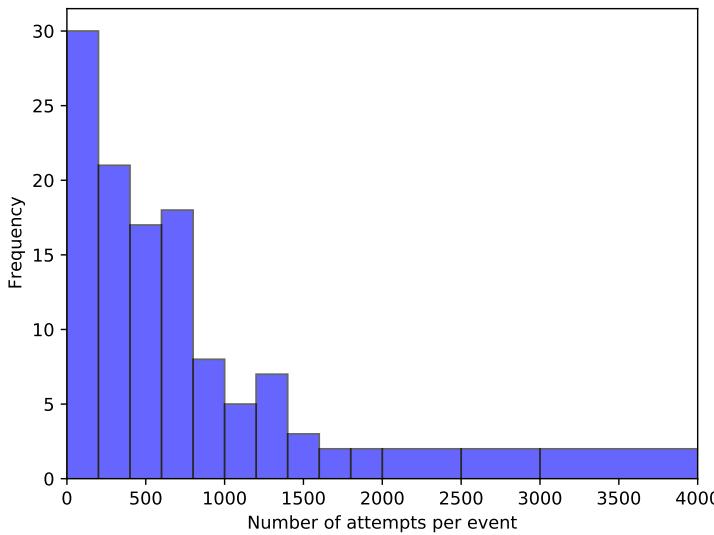
- <sup>5</sup> 1. Propose a test direction  $(\alpha, \phi)$ , by generating  $\phi$  uniformly in the interval  $[0, 2\pi]$ ,  
<sup>6</sup> and  $\alpha$  according to some pre-determined Gaussian distribution, known as the  
<sup>7</sup> Gaussian envelope.
- <sup>8</sup> 2. Given this test direction, calculate where a line following this direction from the  
<sup>9</sup> fibre of interest will hit the PSUP on the other side of the detector. Find the 3  
<sup>10</sup> closest PMTs to that point.
- <sup>11</sup> 3. From those PMTs, obtain their relative intensity values from the beam profile  
<sup>12</sup> mapping, and perform an interpolation based on how close each PMT is to the  
<sup>13</sup> PSUP intersection point. This gives an interpolated relative intensity value for  
<sup>14</sup> this test direction.
- <sup>15</sup> 4. Because we are sampling using the angular coordinates  $(\alpha, \phi)$ , differential area  
<sup>16</sup> elements over this space of directions do not have the same size. We can correct  
<sup>17</sup> for this fact by multiplying our interpolated relative intensity by  $\sin \alpha$ , which  
<sup>18</sup> corresponds to the Jacobian of the direction-space.
- <sup>19</sup> 5. Calculate the value for the Gaussian envelope along this test direction.
- <sup>20</sup> 6. Throw a random number uniformly between 0 and the Gaussian envelope value. If  
<sup>21</sup> the random number is less than the interpolated intensity, then this test direction  
<sup>22</sup> is accepted, and a photon is generated with that direction. Otherwise, we reject  
<sup>23</sup> the direction and try the whole process again.

This generator certainly works, but has a key problem: efficiency. The 1D generator was able to generate a SMELLIE event (that is, to fully specify the starting parameters of all the photons emitted from a fibre) at a speed of  $\sim 1$  ms. However, the 2D generator specified here could take upwards of  $\sim 50$  s *per event* to generate. Because a typical SMELLIE analysis requires simulating many millions of events, the CPU time taken to perform this quickly became unfeasible. Fixing this generator speed problem was a high priority for the SMELLIE analysis.

### 4.1.2 The new generator

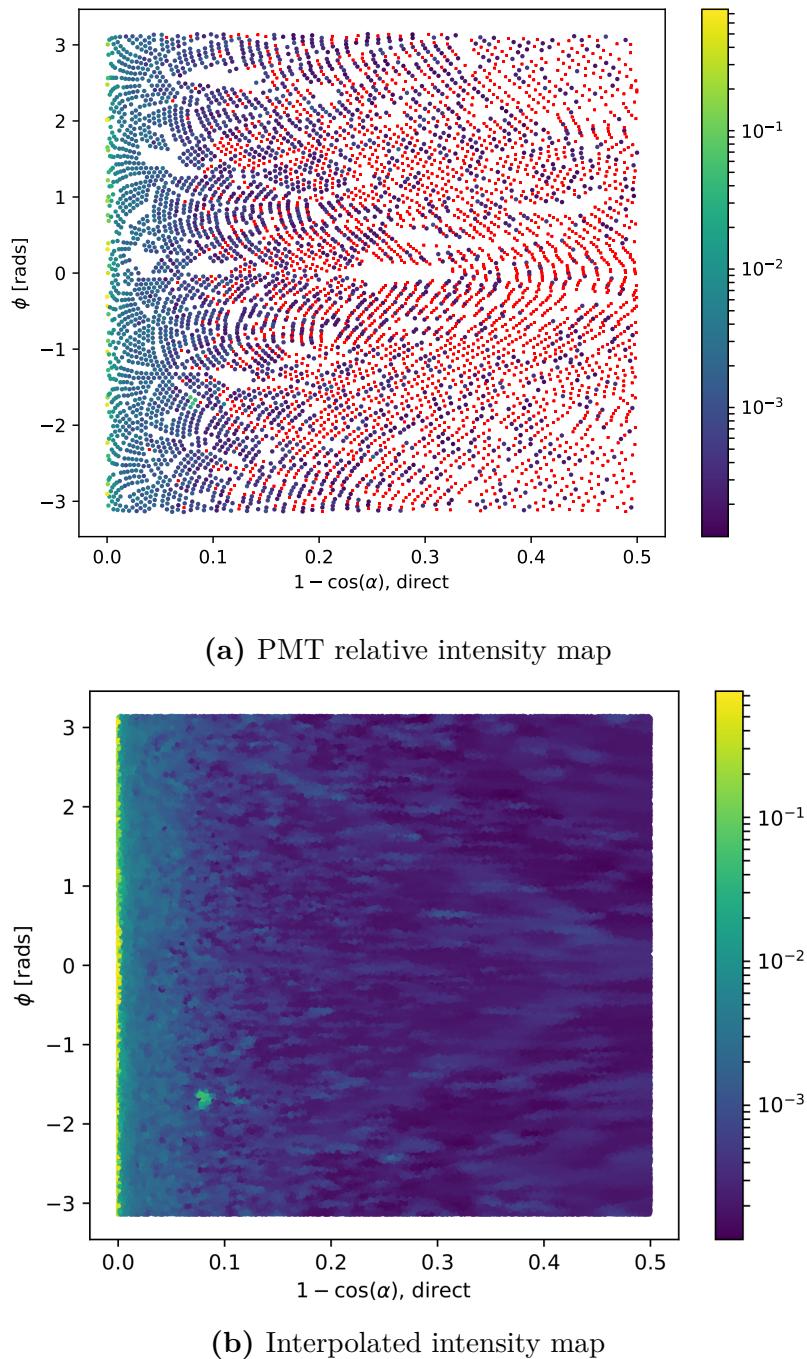
On careful inspection of the existing 2D generator, the main reason for the slowness of the algorithm is the use of a rejection approach. Even with use of the Gaussian envelope, which was included to help with speed, the vast majority of proposed directions are never selected. Figure 4.2 shows a histogram of number of attempts per event it took for a valid direction to be chosen for a representative SMELLIE simulation. Moreover, the calculations needing to be done for every proposed direction are relatively complex, notably trying to find the 3 nearest PMTs to some point on the PSUP.

A new 2D generator was built with these thoughts in mind. Firstly, the rejection method would no longer be used, given its inefficiency. We would also endeavour to try and “pre-calculate” as much as possible before run-time. Starting with the existing PMT relative intensity maps, we plot these in the 2D direction-space ( $1 - \cos \alpha, \phi$ ): see Figure 4.3a. In a toy-MC simulation, 500,000 directions are then thrown uniformly in this 2D space per fibre. For each direction, the same method of obtaining an interpolated intensity value from the nearest PMTs to the corresponding point on the PSUP as from the original 2D generator was performed, the only difference being that these calculations were done well before any actual SMELLIE simulation. Figure 4.3b shows the interpolated intensities obtained for one fibre.

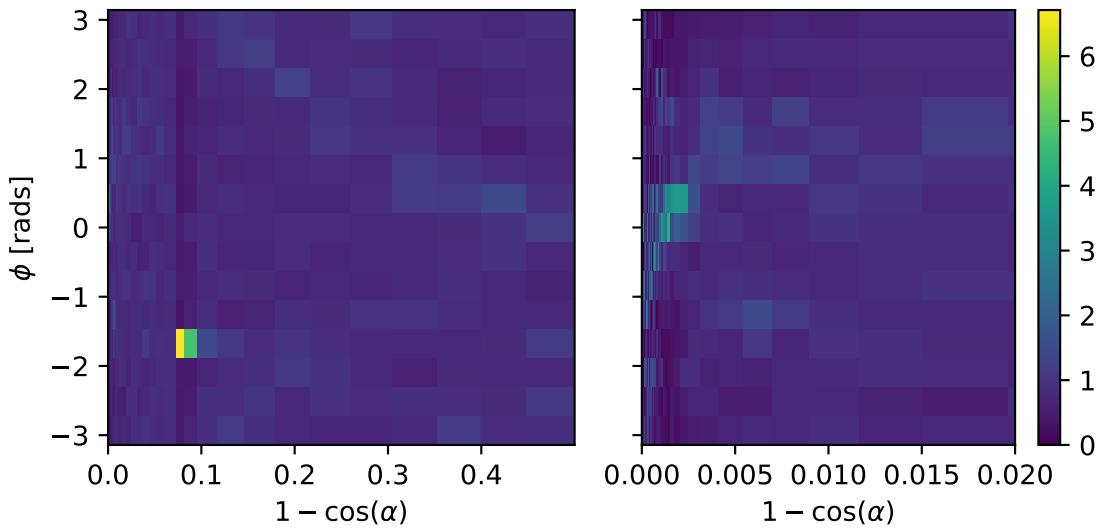


**Fig. 4.2:** Typical distribution of the number of attempts it takes for the existing 2D generator before the test direction gets accepted, per event.

Following this, the sampled intensities were then binned into a 2D histogram, where the bin value corresponds to the sum of all intensities for all directions found within this bin. Choosing a sensible binning procedure is important: too few bins, and necessary information about the shape of the beam is lost, whilst too many bins can oversample the data and capture statistical artefacts in the sampling process instead of just the beam profile. As a balance, 15 bins were chosen along the  $\phi$  direction, and 60 in  $r = 1 - \cos \alpha$ . This was chosen to ensure that a reasonable number of PMTs were located within each bin, lessening the impact of any statistical fluctuations. Although the bins in  $\phi$  were chosen to have uniform width, this was decided to be not the case for the other axis, as there is far more important information near  $r = 0$  (the beamspot). Instead, the width of the bins in  $r$  were calculated so that roughly the same total probability was contained in each  $r$ -strip. By consequence, bins near the beamspot typically are of significantly smaller size than ones much further out. This allows us to both capture any rapid changes in intensity near the beamspot, where this



**Fig. 4.3:** The first step in the new method for preparing the new generator. In (a), the relative intensities used for the existing beam profile of fibre labelled FS055 are shown for each PMT, the position on the plot indicating the location of that PMT in the fibre coordinates. The colour indicates the relative intensity; PMTs marked red have an intensity of zero. Figure (b) shows the result of throwing 500,000 directions uniformly over this 2D space, the intensity of each point given by interpolating the intensities of nearby PMTs.



**Fig. 4.4:** Histogram of interpolated intensities within the 2D direction-space. The left view shows the full histogram; the right is a zoomed-in version near the beamspot. Unlike the binning in  $\phi$ , the bin widths in  $r$  are not at all uniform. Instead, they have been determined such that the area summed over a given “strip” of bins of constant  $r$  will be the same.

<sup>1</sup> matters greatly, and smooth out the very-low intensities seen at larger polar angles.

<sup>2</sup> One of these histograms can be seen in Figure 4.4: the large change in bin widths as a  
<sup>3</sup> function of  $r$  is clear. One can also see that near the beamspot notable dependence on  
<sup>4</sup> the intensity as a function of  $\phi$ . The mysterious “spot” at  $r = 0.08$ , well out of the  
<sup>5</sup> beamspot, is an indication that the underlying beam profile data being used requires  
<sup>6</sup> improvement: more on this in section 4.2.

<sup>7</sup> The Cumulative Density Function (CDF) of this intensity histogram as a function  
<sup>8</sup> of bin was then produced, where the bins were ordered through a raster-scan: scanning  
<sup>9</sup> first over  $\phi$ , and then  $r$ . The CDF was then normalised to 1 so that it was well-defined.

<sup>10</sup> It is this CDF object that is then loaded in and sampled from during event generation.  
<sup>11</sup> To do this, an “inverse-CDF” approach was used, which has the major benefit over  
<sup>12</sup> rejection sampling of always producing a valid direction for every sample made. The  
<sup>13</sup> algorithm works as follows:

- <sup>14</sup> 1. Throw a random number uniformly in  $[0, 1]$ .

---

4.2 Improving the beam profiles

51

2. Perform a binary search to find the bin that has the largest CDF value below  
this random number. 1
3. Look at the bin edges in  $\phi$  of this selected bin: use linear interpolation of the  
random number to obtain a  $\phi$  value located between these two  $\phi$ -values. 2 3
4. Look at the selected bin's  $r$ -bin edges, and select a value of  $r$  by throwing a  
second random number uniformly between the two edges. Convert this  $r$  into a  
polar angle  $\alpha$ . 4 5 6 7
5. The photon's direction is defined by the  $(\alpha, \phi)$  chosen by this process. 8

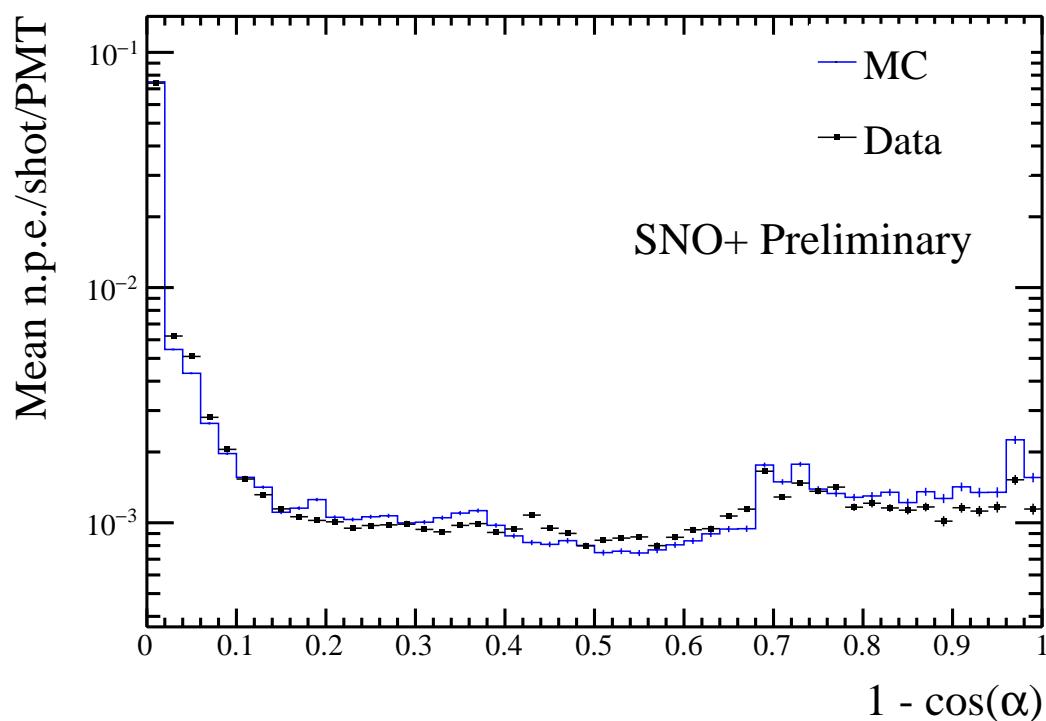
Because of the relative simplicity of this algorithm compared to the previous 2D generator, the speed improvement was very large: generation now took  $\sim 1$  ms per SMELLIE event, a speed improvement of nearly 50,000. Event generation became as fast as it was when the 1D generator was being used. Furthermore, because of the approach taken, this major speed improvement comes at no sacrifice in accuracy. Figure 4.5 shows a comparison of the average number of photoelectrons (npe) per event per PMT between water-phase SMELLIE data and simulations with both the old and new 2D generator. One can see clearly that both generators are as accurate as one another. Note that this plot uses the updated beam profiles as explained in the next section. 9 10 11 12 13 14 15 16 17 18

---

**4.2 Improving the beam profiles**

19

Even with the new 2D profile generator, a problem remains: the simulation fails to reasonably recreate data, and much of this appears to be because of the poor beam profile data being used. The curious “spot” for one of the fibres was already noted in the previous section that doesn't seem to be physical, and more broadly at large angles 20 21 22 23



**Fig. 4.5:** Comparison of water-phase data to MC generated using both the old and new 2D beam profile generator approaches, with the updated beam profiles. Both versions of the generator are consistent with one another, but the new generator is many times faster.

---

4.2 Improving the beam profiles

## 53

for all the fibres there are large swathes of PMTs with an intensity of zero, providing  
 little useful information about the beam shape. It was shown in [32] that with the  
 old 2D generator, the systematic uncertainty on the beam profiles was the dominant  
 source of error in the main SMELLIE analysis. To help improve this situation, it was  
 decided to update the existing beam profiles.

These old beam profiles were originally determined by looking at SMELLIE data  
 taken during the water-phase. Specifically, a “medium”-intensity subrun with one  
 of the lasers firing at a wavelength of 495 nm, was chosen for each fibre. “Medium”-  
 intensity corresponds to firing the relevant laser at a set intensity determined during  
 an earlier commissioning process, for which the maximum occupancy of PMT hits at  
 that intensity, i.e. the proportion of hits per event, corresponded to roughly 80%. This  
 value was chosen as it allowed for high statistics in a relatively short run-time, but not  
 so intense that the occupancy of any given PMT in the beamspot was 100%. Because  
 Rayleigh scattering is strongly-dependent on wavelength, the long wavelength of light  
 was chosen so that impacts from this scattering were small in the data.

SNO+ PMTs are unable to distinguish the exact number of photoelectrons being  
 generated. One is typically only able to know if a PMT has been triggered at all, by  
 any number of photoelectrons. As a result, the occupancy of a PMT over a number  
 of SMELLIE events,  $o$ , is a biased estimator of the mean number of photoelectrons  
 generated,  $\mu$ . Assuming the number of photoelectrons generated in a given event  
 follows Poisson statistics, the probability of generating  $k$  photoelectrons is:

$$P(k|\mu) = \frac{\mu^k e^{-\mu}}{k!}. \quad (4.1)$$

The probability of observing a “hit” in a given PMT corresponds to generating at  
 least one photoelectron:

$$P(\text{hit}|\mu) = P(k \geq 1|\mu) = 1 - P(k = 0|\mu) = 1 - e^{-\mu}, \quad (4.2)$$

which implies after rearrangement that one can determine the mean number of photoelectrons per event from the occupancy by:

$$\mu = \ln(1 - o). \quad (4.3)$$

This is the reason why we want to avoid PMTs with occupancies of 100%: they preclude one's ability to convert into a value for  $\mu$  by looking at occupancy alone. We call this conversion from occupancy into npe the “multi-hit correction”. The impact of this correction is typically small for most PMTs, but can become very significant in a fibre's beamspot.

Once the npe mapping from data was obtained, a correction was then made for the detector's optics: even ignoring a fibre's beam profile, we still expect certain PMTs to be illuminated more than others because of e.g. reflections off the AV, or the solid angle subtended by the PMT bucket opening. For each fibre, a simulation was made where the beam profile was set as isotropic, and the corresponding npe mapping obtained: this map held information about the detector optics only. The beam profile mapping was then derived by simply dividing each fibre's npe mapping from data to its associated isotropic MC npe map. It is these maps that were first used in section 4.1.2.

### 4.2.1 Combining beam profile datasets

Fortunately, much more SMELLIE data was taken during the water-phase than was used for the original beam-profiling analysis. This additional data can be combined with that which was already used to far better constrain the beam profiles. In particular, given the existing assumption that scattering effects are minimal above wavelengths

## 4.2 Improving the beam profiles

55

Run Number	Run Type	Comments
114,018	All PQ lasers; SuperK laser in 400–500 nm range	Only PQ495 laser and SuperK at 495 nm is used
114,023	SuperK laser in 500–600 nm range	Part 1 of this wavelength range; crash occurred on last subrun, so that subrun is ignored
114,034	SuperK laser in 500–600 nm range	Part 2 of this wavelength range

**Table 4.1:** Water-phase runs used for new beam profiling.

of  $\sim 490$  nm, all data taken with wavelengths above this can also be used. The specific runs (and associated comments about their specifics) are described in Table 4.1. Because high-intensity runs require a different analysis approach (PMTs with high occupancies must use charge, not occupancy, to estimate npe), for this analysis we only considered subruns that used low or medium intensity set-points.

For each subrun  $j$  of data per fibre, we look only at PMT hits for each PMT  $i$  that has been identified as “good” for that subrun<sup>1</sup>,  $i \in G_j$ .  $G_j$  here represents the set of good PMTs in subrun  $j$ . In particular, a “good” PMT must have valid electronic and timing calibrations, be at high voltage and masked into the detector’s trigger system for that subrun. In addition, an angular cut of  $\alpha < 60^\circ$  was made to remove PMTs that are well outside any reasonable beam direction. To isolates the hits arriving directly from the fibre without reflecting, scattering, or being noise, a time cut was also made. Because what matters is the time relative to emission from the fibre, and the expected time-of-flight from fibre to different PMTs varies, a quantity known as the time residual was used. Starting with the calibrated hit time of a given PMT relative to the event’s trigger time,  $t_{hit}$ , the expected time-of-flight  $t_{TOF}$  from the fibre to the PMT was subtracted, estimated with the collaboration’s “Light Path Calculator”. Then, the emission time was also subtracted,  $t_{emm}$ , estimated by looking at the second-earliest value of  $t_{hit} - t_{TOF}$  within the fibre’s central beamspot, defined as the PMTs for which

<sup>1</sup>Strictly speaking, a PMT’s “goodness” is only determined on a run-by-run, not a subrun-by-subrun level, but this has no impact on the analysis.

<sup>1</sup>  $\alpha < 3^\circ$ . It was found that a “loose” time residual cut of  $t_{res} \in [-10, +12]\text{ns}$  was  
<sup>2</sup> sufficient to remove the vast majority of non-direct light with little signal sacrifice. In  
<sup>3</sup> the situation where a subrun with intensity was very small, it would not regularly have  
<sup>4</sup> at least two hits in the beamspot, and so the time residuals calculated would not be  
<sup>5</sup> valid for many events. To avoid this situation, a cut was made on any subruns with  
<sup>6</sup> mean intensities below 9 within their beamspot. This value was chosen as it would  
<sup>7</sup> mean a  $2\sigma$  fluctuation downwards of  $2 \cdot \sqrt{9} = 2 \cdot 3 = 6 \text{ npe}$  would still have more than  
<sup>8</sup> the 2 hits necessary for timing reconstruction. One fibre, FS207, has no data subruns  
<sup>9</sup> that satisfy this condition, and as such will have to be dealt with separately. For the  
<sup>10</sup> time being, this fibre was ignored.

<sup>11</sup> Extracting the underlying beam profiles from these data required some careful  
<sup>12</sup> thought, especially because different subruns could have different intensities. Consider-  
<sup>13</sup> ing a PMT  $i$  in subrun  $j$ , the mean number of photoelectrons generated per event in  
<sup>14</sup> that PMT for that subrun,  $\mu_{ij}$  can be decomposed as follows:

$$\mu_{ij} = I_j k_i = I_j b_i f_i. \quad (4.4)$$

<sup>15</sup>  $I_j$  is the intensity of the subrun, i.e. the mean number of photons generated from the  
<sup>16</sup> fibre in that subrun per event.  $k_i$  is the probability that a given photon generated at  
<sup>17</sup> the fibre source ends up generating a photoelectron in PMT  $i$ . This itself can be further  
<sup>18</sup> split into two components:  $b_i$ , the probability that a given photon at the fibre source  
<sup>19</sup> points in the direction of PMT  $i$ ; and  $f_i$ , the probability that a given correctly-pointed  
<sup>20</sup> photon actually makes it to the PMT and successfully generates a photoelectron. It is  
<sup>21</sup>  $b_i$  that is the actual beam profile we would like to measure.  
<sup>22</sup>

<sup>23</sup> Letting  $p_{ij}$  be the probability of observing a hit for a given event on a given  
<sup>24</sup> PMT, the probability of observing  $m_{ij}$  hits out of  $N_j$  events in the subrun will be

4.2 Improving the beam profiles

57

binomially-distributed:

$$P(m_{ij}|\mu_{ij}) = L(\mu_{ij}|m_{ij}) = \binom{N_j}{m_{ij}} p_{ij}^{m_{ij}} (1-p_{ij})^{N_j-m_{ij}} = \binom{N_j}{m_{ij}} (1-e^{-\mu_{ij}})^{m_{ij}} e^{-\mu_{ij}(N_j-m_{ij})}. \quad (4.5)$$

Here we have used equation 4.2, and noted that this probability distribution in  $m$  can be re-framed as a likelihood function for the parameter  $\mu_{ij}$ . Considering only a single subrun of data, the maximum likelihood estimate of the parameter  $\mu_{ij}$  can be shown to be:

$$\langle \mu_{ij} \rangle = -\ln \left( 1 - \frac{m_{ij}}{N_j} \right) = \ln (1 - o_{ij}) \quad (m_{ij} \neq N_j), \quad (4.6)$$

where  $o_{ij}$  is just the occupancy of PMT  $i$  in subrun  $j$ . This is just the multi-hit correction formula seen in equation 4.3, which makes sense.

When looking at multiple subruns for the same fibre, the total likelihood function for a given PMT when considering all the data for a given fibre will be the product of the likelihoods from each dataset,

$$L(\{I_j\}, k_i | \{m_{ij}\}) = \prod_j L(I_j, k_i | m_{ij}) = \prod_j \binom{N_j}{m_{ij}} (1 - e^{-I_j k_i})^{m_{ij}} e^{-I_j k_i (N_j - m_{ij})}. \quad (4.7)$$

This leads to a log-likelihood distribution of

$$\mathcal{L}(\{I_j\}, k_i | \{m_{ij}\}) = \sum_j \left[ \ln \left( \binom{N_j}{m_{ij}} \right) + m_{ij} \ln \left( 1 - e^{-I_j k_i} \right) - I_j k_i (N_j - m_{ij}) \right]. \quad (4.8)$$

Formally, one could combine the likelihoods of all the PMTs together, and by looking at the maximum likelihood estimates for each of the parameters measure the parameter values this way. However, the set of equations one obtains through this approach quickly become analytically intractable, because the PMTs are coupled by the intensity values  $I_j$ . Even a direct numerical approach would be liable to fail: for a given fibre

<sup>1</sup> there can be dozens of subruns, and many thousands of PMTs of relevance, so the  
<sup>2</sup> dimensionality of the system of equations would be far too large.

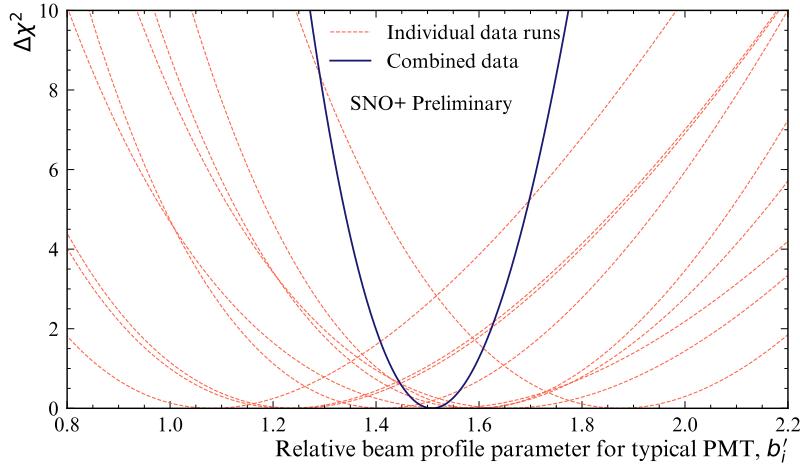
<sup>3</sup> Because of this, a different approach was taken. It is expected that in a subrun the  
<sup>4</sup> total npe, summed over all good PMTs, should be proportional to the intensity value  
<sup>5</sup>  $I_j$ . One must be careful about this construction — different subruns can have different  
<sup>6</sup> sets of good PMTs, so two subruns with identical  $I_j$  values could have a larger summed  
<sup>7</sup> npe merely because more PMTs were good in that subrun. To counter-act this effect,  
<sup>8</sup> only PMTs that were classified as good in *all* subruns being analysed for that fibre  
<sup>9</sup> would be used for the npe summation. In other words, we use data from PMT  $i$  for  
<sup>10</sup> summing only if:

$$\text{11} \quad i \in \mathcal{I} = \bigcap_j G_j. \quad (4.9)$$

<sup>12</sup> We can then define the summed npe for a given subrun as  $S_j = \sum_{i \in \mathcal{I}} \text{npe}_{ij}$ , and assert  
<sup>13</sup> that  $I_j = cS_j$ . By finding a value proportional to  $I_j$ , there is now enough information  
<sup>14</sup> to maximise the log-likelihood  $\mathcal{L}(k_i | \{m_{ij}\}, \{I_j\})$  with respect to  $k_i$  for each PMT  
<sup>15</sup> independently, and hence obtain estimates for these  $k_i$  parameters.

<sup>16</sup> Of course, what is actually wanted are the underlying  $b_i$  values, not  $k_i$ . This is  
<sup>17</sup> where isotropic simulations come in. For each run of data used, a matching isotropic  
<sup>18</sup> MC was produced. As an example, a simulation for run 114,023 contained 200,000  
<sup>19</sup> events for each fibre using an isotropic beam profile, over the full wavelength range  
<sup>20</sup> considered in this run, 500–600 nm, using the same run conditions as in data (which  
<sup>21</sup> PMTs were at high voltage, etc.).

<sup>22</sup> For each isotropic MC run, both  $I_j^{MC}$  and  $k_i^{MC}$  were calculated via the method  
<sup>23</sup> described above. Because the simulations were isotropic, the underlying value for  $b_i$   
<sup>24</sup> was constant across all the PMTs, and so  $ak_i^{MC} = f_i$ . By doing some rearranging of



**Fig. 4.6:** Plot of  $\Delta\chi^2 \simeq X_i$ , twice the negative log-likelihood ratio, for both single subruns of a typical PMT, and when all relevant subruns are combined.

equation 4.4, we find that:

$$\mu_{ij} = I_j b_i f_i = c S_j b_i a k_i^{MC} = (acb_i) S_j k_i^{MC}. \quad (4.10)$$

As a result of this, given the set  $\{S_j\}$  and  $k_i^{MC}$ , one can maximise the log-likelihood  $\mathcal{L}$  with respect to  $b'_i = acb_i$  numerically, to obtain the maximum likelihood estimate of  $b'_i$ . Because  $a$  and  $c$  were global constants of proportionality, they would become irrelevant as soon as the beam profile was normalised in the CDF-creation process outlined in 4.1.2.

Figure 4.6 shows the shape of this log-likelihood distribution for a particular PMT when considering fibre FS007's beam profile. One can see how individual subruns provide much more information when combined than if one looked at a single subrun alone.

Another benefit of using this log-likelihood approach is that the resulting distribution's shape can be used for uncertainty estimation. In almost all cases, Wilks Theorem [48] allows us to produce  $1\sigma$  confidence intervals about the maximum likeli-

hood estimate for  $b'_i$ ,  $\langle b'_i \rangle$ , because

$$X(b'_i) = -2 [\mathcal{L}(b'_i) - \mathcal{L}(\langle b'_i \rangle)]$$

<sup>1</sup> approximates a  $\chi^2$ -distribution. As a result, the error bounds on our parameter estimate  
<sup>2</sup> are given by when  $X = 1$ . The fact that the shape of  $X$  can be well-approximated by  
<sup>3</sup> a quadratic in the region near  $X = 0$  indicates the validity of Wilks' Theorem being  
<sup>4</sup> used here.

<sup>5</sup> Only a couple of exceptions to this approach of parameter estimation are possible.  
<sup>6</sup> In the case where  $m_{ij} = N_j$ , i.e. a PMT has 100% occupancy, no maximum likelihood  
<sup>7</sup> estimate exists: we need not worry about this, as subruns where this occurs have not  
<sup>8</sup> been used. On the other end, however, there are some PMTs for certain fibres where  
<sup>9</sup> after all subruns of data have been included, there remains no hits. In this scenario,  
<sup>10</sup> one can show that the log-likelihood becomes linear in the beam profile parameter:

$$\mathcal{L}(b'_i | \{m_{ij} = 0\}) = b'_i k_i^{MC} \cdot \sum_j [I_j N_j]. \quad (4.11)$$

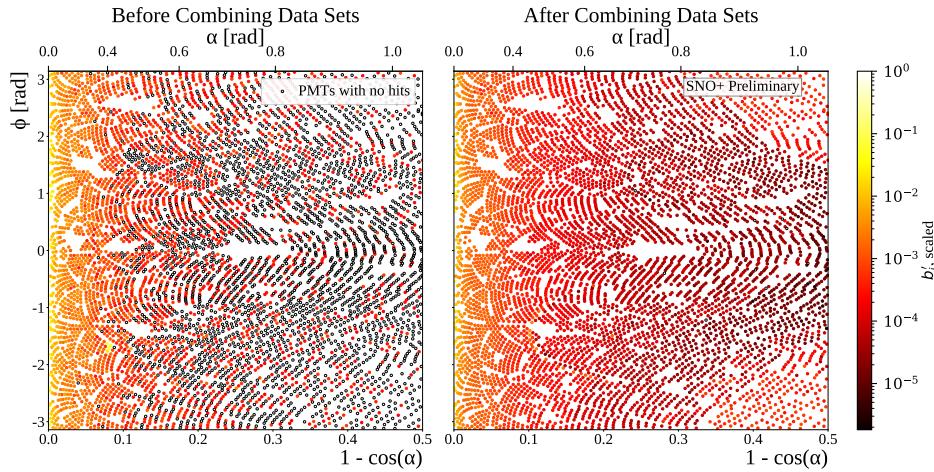
<sup>12</sup> This scenario is very much reminiscent of rare-decay searches, and a similar approach  
<sup>13</sup> can be used. A  $1\sigma$  upper limit on the possible value for  $b'_i$  can be analytically-calculated  
<sup>14</sup> to be:

$$b'_{i,ulim} = -\frac{k_i^{MC} \sum_j [I_j N_j]}{\ln [1 - \text{erf}(1/\sqrt{2})]}, \quad (4.12)$$

<sup>16</sup> where  $\text{erf}(x)$  is the error function. [18 pages for above two sections]

### <sup>17</sup> 4.2.2 Results & Discussion

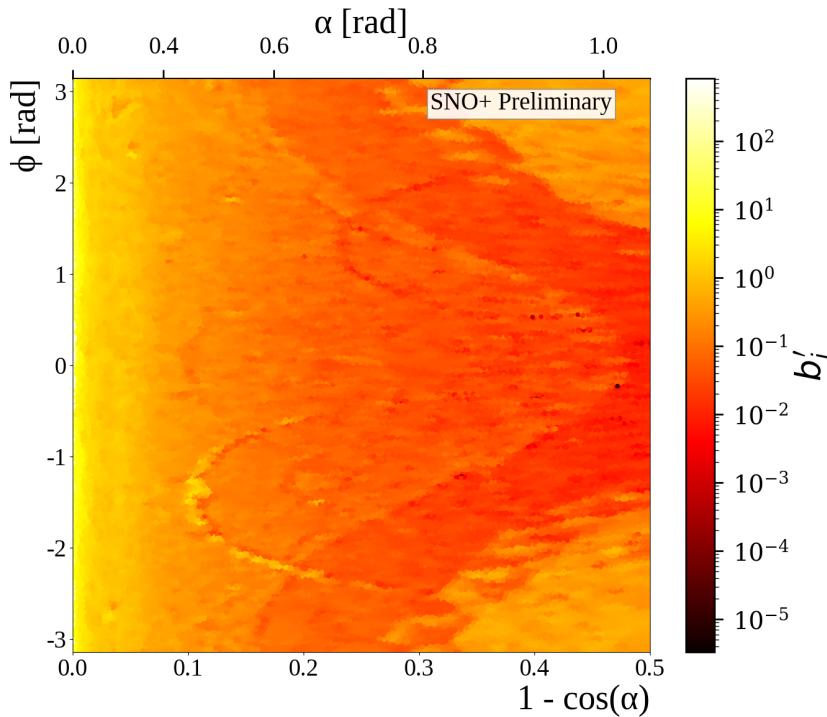
<sup>18</sup> WARNING: contents of this subsection will be gutted, focusing merely on impact of  
<sup>19</sup> combining data sets. Details about discrepancies will be covered in the next section.



**Fig. 4.7:** Comparison between old and updated beam profiles for fibre FS055, after combining multiple data sets. Once again, the relative intensities ( $b'_i$ ) for each PMT are given by the colour of each point, the position of each plotted in the 2D  $(r, \phi)$ -space. The relative intensities have been both scaled here so that the largest value equals 1. Hollowed-out points are PMTs that, even after all relevant subruns have been combined, have no PMT hits.

[Will just be 1 page] Figure 4.7 shows the impact of using additional subruns of data on a typical beam profile. One can clearly see the great reduction in the number of PMTs with no hits in data. That many more data sets were included allowed for the major increase in dynamic range available for measuring these  $b'_i$  values. One can also note that by including additional data the curious spot that was seen in the old beam profile our at  $r \approx 0.08$  has gone, further indicating that it was an artefact of that single data set.

Further details can be gathered from the interpolated intensity maps, one of which can be seen in figure 4.8. There are two curious stand-out features that can be seen here: firstly, there are multiple distinct parabolic arcs. These correspond to the shadows of the ropes that hold up/down the AV. More precisely, they are the mismodelling of those shadows — if the shadows were in the right place in the isotropic MC, then they would correctly cancel out any decreased intensity seen in the data of shadowed PMTs. These shadows could be mismodelled either because the positions of the ropes in the MC are in the wrong place, or the fibre's emission position is wrong. Note that any



**Fig. 4.8:** Interpolated intensity map for the new updated beam profile of fibre FS055. The misalignment of rope shadows and AV effects, can both be seen.

1 mismodelling of the fibre's nominal emission direction has no impact on this shadowing  
 2 problem, as changing that direction merely causes a change of basis in the  $(r, \phi)$ -space.  
 3 The latter possibility of incorrect fibre positions are more likely, and in fact these arcs  
 4 in the beam profiles could be used as an effective way to correct for this problem.

5 The second distinctive feature of this intensity map is the large band of lower  
 6 intensity varying between  $r \approx 0.2 - 0.5$ , followed by larger intensity out at large  
 7  $r$  values. This feature comes from light reflecting off the AV surface, or internally-  
 8 reflecting. The reason for this band's functional dependence on  $\phi$  is that this particular  
 9 fibre, FS055, has a nominal fibre direction  $\sim 10^\circ$  from pointing radially-towards the  
 10 detector's centre. This feature appears in the updated beam profiles of all fibres, but  
 11 its shape depends on the particular fibre's direction — for fibres pointing directly  
 12 towards the detector's centre, there is little  $\phi$ -dependence observed. Like the ropes,

---

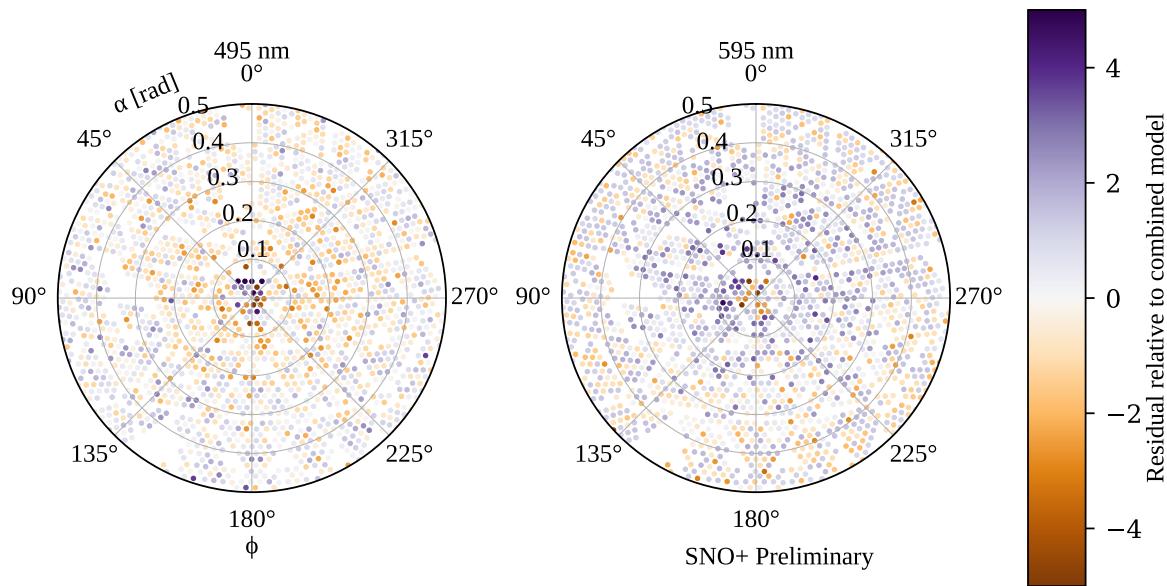
4.2 Improving the beam profiles

## 63

this feature must come from some form of mismodelling of the optics of the AV. A de-facto shadowing of PMTs in line with tangents from the AV surface which intersect the fibre position is to be expected. One also expects PMTs at polar angles larger than this to have their observed intensities boosted from reflected light off the AV. However, the discontinuities seen in the beam profiles indicate that for whatever reason this effect has been over-emphasised in the simulation.

There is a further phenomenon that can be seen, by comparing beam profile values obtained from a single subrun to the updated combined beam profile. This can be done by calculating the residuals corresponding to the single subrun, relative to the combined data set. The residual is negative if the combined data sets have a  $b'_i$  below the equivalent for a given single subrun; that is, the combined model underestimates this subrun for that PMT.

This information was plotted for two different subruns from the same fibre, seen in figure 4.9. One subrun was the same one used by Esther Turner for the original 2D beam profiling, with a wavelength of 495 nm; the latter was at the longer wavelength of 595 nm. For both subruns, most PMTs are seen to have intensities well-modelled by the combined model. However, there appears to be a significant amount of mismodelling within the beamspot. There also appears to be some systematic shift between data and model at somewhat larger polar angles. Moreover, this mismodelling seems not to be merely random, but a function of wavelength: at shorter wavelengths the beamspot tends towards being overestimated and then underestimated at larger values of  $\alpha$ . At longer wavelengths, the beamspot becomes underestimated, with larger angles getting overestimated. This indicates that there appears to be a wavelength-dependence on the beam profiles, contradicting one of the main assumptions which we used to combine the water-phase data in the first place! All three of these features — rope shadows, AV reflections, and wavelength dependence — add systematic uncertainty to the beam



**Fig. 4.9:** Residuals from subruns at two different wavelengths, both compared to the combined beam profile model for fibre FS055. A negative sign, and hence bluer colours, indicate that the combined model underestimates the observed intensity for that particular subrun. Values with a magnitude beyond 5 are shown capped at this maximal value for the purposes of this plot. These PMTs are plotted in the polar fibre coordinates  $(\alpha, \phi)$ .

1 profiles, beyond the statistical uncertainty as measured by the width of the likelihood  
 2 distribution. Certainly if one wanted to further improve the uncertainties in the beam  
 3 profiles, tackling these challenges would be key.

### 4 4.3 Comparisons between Data and Simulation

- 5 • This focuses on disagreements noticed between data and MC even after the new  
 6 generator and beam profiles have been used.
- 7 • Important to mention that none of these are necessarily game-ending, they are  
 8 just systematics that may or may not be substantial in a given analysis with  
 9 SMELLIE..

### 4.3.1 Forward Hemisphere Discrepancies

- The continued disagreement between data and MC when it comes to measured npe in various parts of the “forward” hemisphere. This includes:
  - The central beamspot,
  - The TIR region,
  - Rope shadows,
  - A noticeable wavelength-dependence.
- This is pretty much most of the contents of Section 4.2.2.

[4 pages]

### 4.3.2 Emission Time Discrepancies

- For certain lasers, a strong mismatch in the observed hit time residuals for prompt light.
- A mysterious +18 ns bump seen for the PQ495 laser.
- A trigger jitter in the SuperK laser.

[3 pages]

### 4.3.3 Backward Hemisphere Discrepancies

- The observed distribution of hits vs time and angle in MC does not match data in a number of ways for PMTs near the fibre emission point.
- Includes the outer-water scattering length, rope reflections, and investigations into whether certain modifications to the optics could plausibly fix things (so far, no).

<sup>1</sup> [6 pages]

<sup>2</sup> [32 PAGES TOTAL]

# Chapter 5

## Analysis of SMELLIE Data in the Scintillator Phase

This chapter contains two sets of analyses: measurements of the extinction lengths of the scintillator as a function of wavelength and time, as well as monitoring the Rayleigh scattering length over time.

### 5.1 Extinction Length Analysis

#### 5.1.1 Motivation

- Explain motivating observations for this analysis: a substantial discrepancy between MC and data seen in the radial profiles of nhitsCleaned for  $^{210}\text{Po}$  after the PPO top-up campaign. Performed by Serena.
- Hypothesis of a shortening of the absorption/scattering length proposed, further strengthened by Ben Tam’s ex-situ absorption measurements with scintillator taken from the detector, as well as knowledge about a likely “cooking” of PPO during the PPO-fill.

- Describe the provisional optics model decided on based on these measurements, which includes an additional non-re-emitting component of the scintillator.
- As a further cross-check, SMELLIE should be sensitive to changes in the overall extinction length of the scintillator, especially for short extinction lengths relative to the size of the detector.
- More straightforward in measuring extinction length compared to scattering length — no need to distinguish between scintillator re-emission and scattering.
- Further uses: can be used to monitor the extinction length over time!

[4 pages]

### 5.1.2 Analysis Approach

#### The 2-PMT Case

- Outline theoretical approach for how one could measure the extinction length of scintillator through a comparison of SMELLIE data between the scintillator and water phases, in the simplified 1 dimensional case with only 2 PMTs.

[3 pages]

#### Combining Results Between PMTs

- Not doing analysis with just 2 PMTs, of course! Can combine results from multiple PMTs within a beamspot: I explain how here.

[2 pages]

**Corrections Between the Water and Scintillator Phases**

- Note the complications that we have to deal with. Namely, the differing refractive indices of the media bending the beamspot differently in the phases, as well as the method used to estimate  $t_{\text{emm}}$ .
- Explain how we deal with these, the former through MC simulation.

[2 pages]

6

**5.1.3 Validation of the Analysis in Simulation**

- Show results of this approach being used to measure the extinction length in simulation. How well does it do?

[3 pages]

10

**5.1.4 Results in Data**

- Describe the data used in this analysis, both water and scintillator, which can be shown in a table.
- Show examples of analysis of data in action for 375 nm data: typical  $t_{\text{res}}$  distributions of backscattered and beamspot PMTs; calculation of that particular extinction length measurement, followed by the graph for extinction length in 375 nm over all fibres and time periods.
- Discuss what results can be seen in this plot: consistency between fibres, the expected change as a function of PPO concentration, and stability of the extinction length during the main 2.2 g/L scintillator phase.

- Compare results to those made by Ben ex-situ: are they in agreement? If not, what possible systematics could there be? The main one for my analysis is likely to be uncertainties in the simulated beam profile that leak through into the refractive index correction of the beamspot. For the ex-situ analysis, the value of the extinction length obtained is achieved through background subtraction at some long wavelength, and the particular choice of this wavelength can lead to systematic changes in the obtained extinction length.
- Look at results at longer wavelengths: can anything reasonably be said at these longer wavelengths? Why/why not?
- Finally: describe any conclusions that can be reached, in particular whether we can affirm the optics model we use in RAT.

[8 pages]

## 5.2 Scattering Analysis

### 5.2.1 Historical Approaches and the Problem of Systematics

- Comparison to MC is necessary in scattering analysis, compared to merely being needed as a correction factor. This is because of the angular dependence of scattering. As a result, we can be far more susceptible to systematics from poor modelling!
- As a warning, show how Krish's/Esther's approach to the SMELLIE scattering analysis suffers majorly from these systematic effects. Requires describing their analysis approach briefly, and then explaining how the systematics described in Section 4.3 lead to major problems with this approach.

---

## 5.2 Scattering Analysis

**71**

- Motivates the need for either reduced systematics, or an alternative analysis approach that is more robust to them! 1  
2

[2 pages] 3

### 5.2.2 New Methodology 4

#### Signal Region Selection 5

- Propose the new analysis approach: looking at light in the “bad light-path” PMT region. Define what this region is. 6  
7
- Give qualitative argument for why we expect this region to be robust to the beam profile systematics: dominated by the scattered signal as no direct light can make it here, and changes to beam profile should get “smeared out” after scattering. 8  
9  
10  
11
- Show how simulations indicate this should be a region with a very high purity of scattered light, and (assuming all else being equal) robust to beam profile uncertainties. 12  
13  
14
- Confirm robustness of selected PMT region to uncertainties in AV offset and fibre position. 15  
16

[5 pages] 17

#### Measuring the Emission Intensity 18

- Remaining systematics is now in the calculation of an average absolute emission intensity. 19  
20

- <sup>1</sup> • Show how various methods don't work particularly well: whole detector npe,  
<sup>2</sup> beamspot npe, backscattered light npe, "bad light-path" PMTs but at later times.  
<sup>3</sup> Explain why it goes wrong for each method.
- <sup>4</sup> • Look at "beamspot but excepting the central bit": if that works well, then we  
<sup>5</sup> can continue!
- <sup>6</sup> • Otherwise, we'll have to live with measuring relative scattering lengths instead  
<sup>7</sup> of absolute amounts, using the outer water back-scattering as a measure of the  
<sup>8</sup> relative emission intensity.

<sup>9</sup> [4 pages]

<sup>10</sup> **5.2.3 Results**

- <sup>11</sup> • Actually do the proposed analysis on data, versus time and wavelength. Do the  
<sup>12</sup> results seem consistent between fibres? Are they sensible values?

<sup>13</sup> [5 pages] [33 PAGES TOTAL]

# Chapter 6

1

## Solar Oscillation Analysis

2

*Driving out into the Sun*

*Let the ultraviolet cover me up*

*Looking for a Creation Myth*

*Ended up with a pair of black lips*

3

---

*This is the End*

PHOEBE BRIDGERS

Measuring the “solar” neutrino oscillation parameters  $\Delta m_{21}^2$  and  $\theta_{12}$  is one of the principal aims of the SNO+ detector during the scintillator-phase. There are, in fact, two complementary methods of measuring these parameters: the oscillations of anti-neutrinos from terrestrial nuclear reactors, and the oscillations of neutrinos from the Sun.

4

5

6

7

8

This chapter focuses on the latter approach, using  ${}^8\text{B}$  neutrinos coming from the Sun to measure the solar oscillation parameters. An initial background-free study was performed by Javi Caravaca [49], which demonstrated that it was indeed possible to make such a measurement in the detector. The work in this chapter builds on substantially from that analysis. This chapter also draws on the associated reactor anti-

9

10

11

12

13

<sup>1</sup> neutrino analysis built by Iwan Morton-Blake [8], and more broadly from the general  
<sup>2</sup> techniques used in the  $0\nu\beta\beta$  analysis of Tereza Kroupova [47] and Jack Dunger [50].

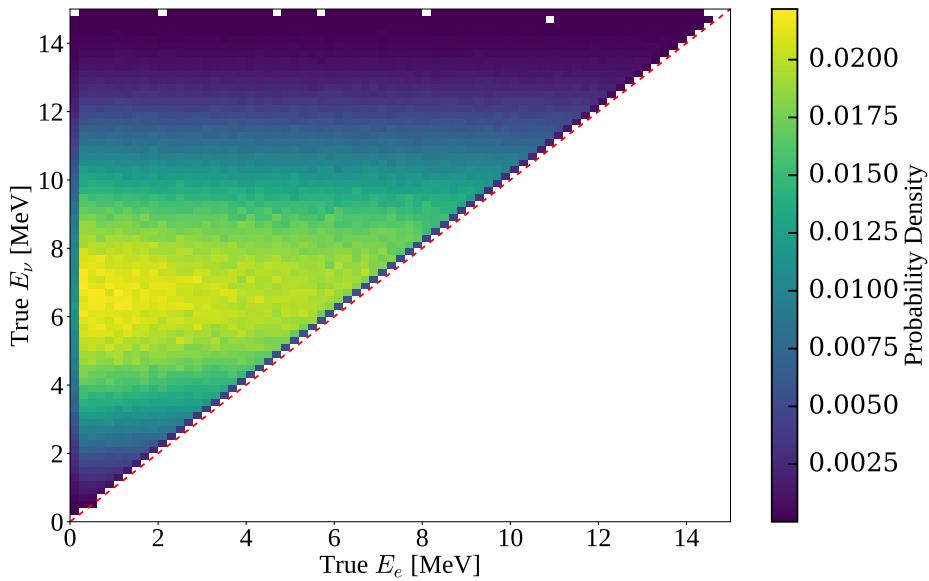
<sup>3</sup> This chapter begins by explaining how it is possible to measure the solar oscillation  
<sup>4</sup> parameters via  $^8\text{B}$  events. Then, the framework used to perform the analysis is then  
<sup>5</sup> explained: that of a *Bayesian Analysis using Markov Chain Monte Carlo techniques*.  
<sup>6</sup> After the method has been described, the dataset upon which the analysis is performed  
<sup>7</sup> is then introduced. The results and associated validation are then given. Given these  
<sup>8</sup> results, a projection is then made for the expected sensitivity to  $\theta_{12}$  as a function of  
<sup>9</sup> livetime.

## <sup>10</sup> 6.1 Analysis Methodology

### <sup>11</sup> 6.1.1 Observational Principle

<sup>12</sup> How can we measure neutrino oscillation parameters via solar neutrinos in the SNO+  
<sup>13</sup> detector? As discussed in Chapter 1, it is possible to detect all flavours of neutrino  
<sup>14</sup> through elastic scattering with electrons in the detector. If this interaction was purely  
<sup>15</sup> neutral-current, then there would be no way of telling the flavour-state of an interacting  
<sup>16</sup> neutrino. However, electron neutrinos are able to interact through an additional  
<sup>17</sup> charged-current mode. This modifies the cross-section for electron neutrinos, and  
<sup>18</sup> means that as the survival probability for electron neutrinos generated from the Sun,  
<sup>19</sup>  $P_{ee}$ , is modified, the interaction probability of neutrinos with the detector will also.

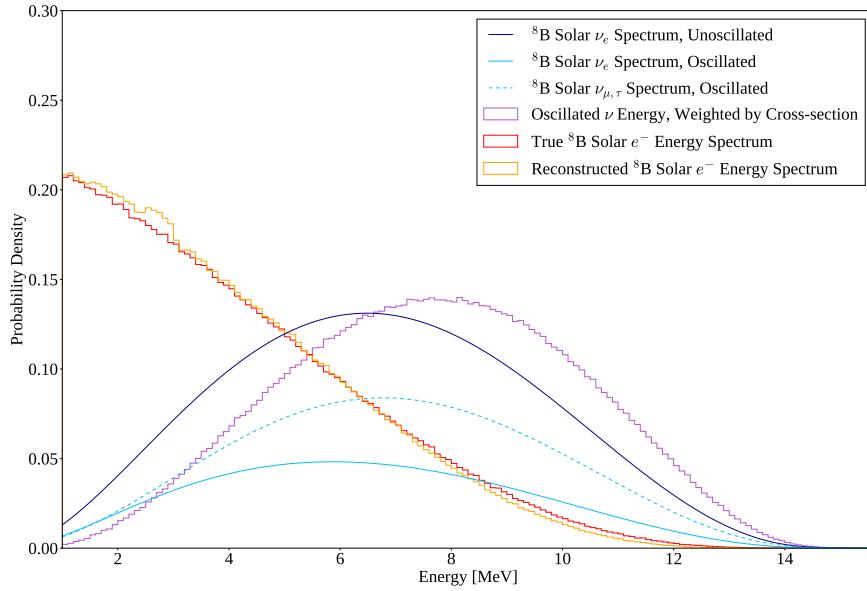
<sup>20</sup> Of course, we do not directly measure neutrino energies in the detector — only  
<sup>21</sup> the associated scattered electron. If there were no correlation between the observed  
<sup>22</sup> electron energy and its associated neutrino, then the only effect of neutrino oscillations  
<sup>23</sup> would be to change the overall observed rate of events due to this process. There  
<sup>24</sup> would be no change in the shape of the event’s energy spectrum, even though neutrino



**Fig. 6.1:** 2D probability distribution comparing the true neutrino energy from a  ${}^8\text{B}$   $\nu_e$  to the true energy of the scattered electron. Also shown in red is the line  $E_\nu = E_e$ .

oscillations are a function of neutrino energy. Fortunately, there is some dependence of  
1  
the neutrino's energy,  $E_\nu$ , on that of the scattered electron,  $E_e$ . This dependence can  
2  
be seen in Fig. 6.1 for  ${}^8\text{B}$  electron neutrinos interacting in SNO+. As can be seen, the  
3  
dependence is weak, and comes mostly from basic energy conservation: If one observes  
4  
a 10 MeV electron event in the detector, it can't reasonably have come from a 5 MeV  
5  
neutrino.  
6

In Fig. 6.2 we can see the impact each physical process has on the energy spectrum  
7  
that we eventually observe. We start with a broad energy distribution of  ${}^8\text{B}$  electron  
8  
neutrinos generated in the Sun. These neutrinos then oscillate their flavour state as  
9  
they propagate to the detector, in an energy-dependent manner. When a (tiny) fraction  
10  
of these neutrinos interact with the electrons in our detector, there is both an energy-  
11  
and flavour-dependence on the cross-section. The scattered electrons gain a kinetic  
12  
energy with some mild dependence on the inciting neutrino's energy, which is then  
13  
measured by the detector to within some energy resolution.  
14



**Fig. 6.2:** The evolution of energy distributions related to  ${}^8\text{B}$  solar neutrino detection. The unoscillated neutrino spectrum is taken from [51]; neutrino oscillations assume oscillation parameters from the current global fit results [52] and  $\nu_e$  survival probabilities calculated via the method described in Section 6.1.8. The latter three distributions were obtained from MC production as described in Section 6.2.1, with the cross-section formula coming from [53].

**Fig. 6.3**

Let us now consider the dependence of  $P_{ee}$  on the individual neutrino oscillation parameters. Recall from Eq. 1.1 that, after considering matter-induced oscillations due to neutrinos passing through the Sun and possibly the Earth,  $P_{ee} = P_{ee}(\tan 2\theta_{12}^M, \sin \theta_{13}^M, \Delta m_{21,M}^2) = P_{ee}(\theta_{12}, \theta_{13}, \Delta m_{12}^2, \Delta m_{13}^2)$ . Fig. 6.3 shows the dependence of each of these four oscillation parameters on  $P_{ee}(E)$ . We can see that in reality only the two parameters  $\Delta m_{21}^2$  and  $\theta_{12}$  have a substantial impact on  $P_{ee}(E)$  and hence the observed electron energy spectrum. Because of this, for this analysis we will only ever vary these two oscillation parameters, and keep  $\theta_{13}$  and  $\Delta m_{13}^2$  at their current global fit values<sup>1</sup> of  $\sin^2 \theta_{13} = 0.0222$  and  $\Delta m_{13}^2 = +2.515 \times 10^{-3} \text{ eV}^2$  [52].

<sup>1</sup>We use the global fit results excluding Super-Kamiokande’s atmospheric data, and assuming normal ordering of the neutrino mass hierarchy. This choice has a tiny impact on the magnitudes of these two fixed parameters, the main impact being the sign of  $\Delta m_{13}^2$ .

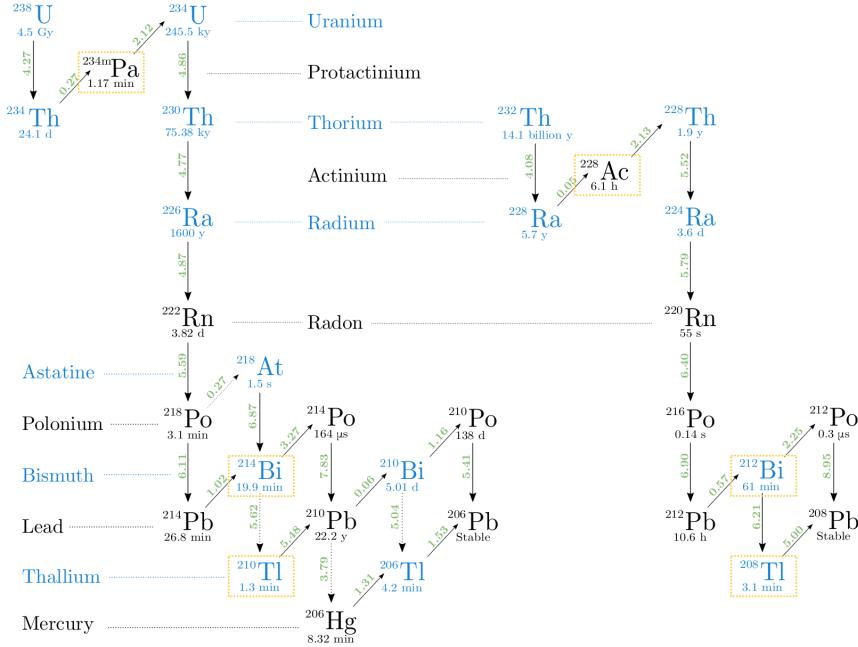
### 6.1.2 Background Processes

Sadly, elastically-scattered electrons from  $^8\text{B}$  neutrinos are not the only events we see in the SNO+ detector during the scintillator phase. There are a number of background processes that our signal must compete against. Below a reconstructed energy of  $\sim 2.5 \text{ MeV}$ , it is known that various backgrounds completely swamp any possible  $^8\text{B}$  signal, and so for this analysis we only consider processes that can generate reconstructed energies of at least  $E_{\min} = 2.5 \text{ MeV}$ . The following subsections explain each of these backgrounds, as well as methods that have been used to mitigate them as much as possible.

#### Internal Uranium- and Thorium-Chain Backgrounds

Although every effort has been made to make the scintillator cocktail that fills SNO+ to be as radio-pure as possible, there inevitably remain trace amounts of the radioactive isotopes that derive from the decay chains of the  $^{238}\text{U}$  and  $^{232}\text{Th}$  isotopes. Fig. 6.4 shows these two decay chains. Fortunately, only a fraction of the radioactive isotopes in these chains actually are capable of generating events in the detector with energies above  $E_{\min}$ : these have been highlighted in Fig. 6.4 in gold.

Of particular note are the decays of  $^{212}\text{Bi}$  and  $^{214}\text{Bi}$ . Both are capable of either  $\alpha$ - or  $\beta$ -decays to Tl or Po isotopes, respectively. For the former, it is the subsequent  $\beta$ -decay of the Tl that can have a reconstructed energy above  $E_{\min}$ . For the latter, the Bi decay is the part of the pair of decays that can lie above  $E_{\min}$ . Although the  $\alpha$ -decays here certainly have Q-values well above 2.5 MeV, the liquid scintillator quenches the observed energy well below  $E_{\min}$ . The so-called “Bi–Po” decays are particularly special because the lifetimes of  $^{212}\text{Po}$  and  $^{214}\text{Po}$  are 300 ns and 164  $\mu\text{s}$ , respectively, which are short enough to allow for highly-effective coincidence tagging.



**Fig. 6.4:** The  $^{238}\text{U}$  and  $^{232}\text{Th}$  decay chains, taken from [47]. Isotopic half-lives are given below their symbol; the Q-values for each decay, in MeV, is given in green. Downward arrows indicate an  $\alpha$ -decay; diagonal arrows indicate  $\beta$ -decay. Isotopes highlighted in gold are potential backgrounds for this solar analysis.

There are two classes of Bi–Po event in the detector: “out-of-window” (OOW) events for which the Bi and Po occur in separate event windows, and “in-window” (IW) events whereby the Bi and Po occur within the same event window. These lead to two distinct strategies for tagging these kinds of events. For out-of-window Bi–Pos, we look for a delayed coincidence of two events. Using the tagging algorithm suggested in [] as a starting point, the chosen procedure was as follows. There must be two events that trigger the detector within  $4\ \mu\text{s}$  of one another, and both have a valid scintFit position reconstruction within  $2\ \text{m}$  of one another. The delayed candidate event must also have at least 100 cleaned PMT hits. This very broad coincidence tagging procedure was designed to ensure that the cut was as *efficient* in tagging (and hence, rejecting) Bi–Pos as possible, whilst negligibly impacting the solar signal. This is in contrast to the cuts chosen by Rafael Hunt-Stokes in [], which try and obtain a highly *pure* sample of Bi–Po tags.

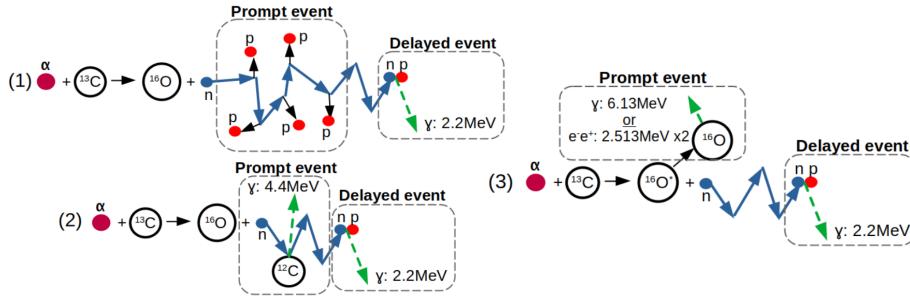
**Fig. 6.5**

Of course, the above delayed coincidence procedure cannot catch any of the in-window Bi–Po events. For these, we use a different approach. Because two decays happened in the same event window, we expect to see two distinct peaks in the event’s time residual spectrum. In order to look for this event topology, a likelihood-ratio classifier was run over events, first developed by Eric Marzec [] and re-coordinated for the  $2.2\text{ g l}^{-1}$  LABPPO scintillator optics by Ziping Ye []. This classifier calculates the likelihood ratio between the null hypothesis of a  $0\nu\beta\beta$  event (a proxy in this analysis for single-site events such as our  ${}^8\text{B}$  signal) and the alternative hypothesis of an in-window Bi–Po event. The more negative the value of the result, `alphabeta212`, the greater the evidence there is for rejecting the null hypothesis of a single-site event. Events with  $\text{alphabeta212} < 0$  were then rejected.

Combining both out-of-window and in-window Bi–Po tagging, the impact on  ${}^{212}\text{Bi–Po}$ ,  ${}^{214}\text{Bi–Po}$ , and  ${}^8\text{B } \nu_e$  events can be seen in Fig. 6.5. We consider here only events that pass all other cuts used in this analysis: see Section 6.2.2 for the specifics of the cuts used. Because of the different lifetimes of the decays,  ${}^{214}\text{Bi–Po}$  decays predominantly fall out-of-window whilst  ${}^{212}\text{Bi–Po}$  events are typically in-window. This explains why the out-of-window tagging is substantially better at cutting  ${}^{214}\text{Bi–Po}$  decays, whereas the in-window tagging far better tags  ${}^{212}\text{Bi–Po}$  decays. Overall, within the analysis region of interest (ROI), the two combined cuts are able to tag TODO% of  ${}^{214}\text{Bi–Po}$  triggered events, TODO% of  ${}^{212}\text{Bi–Po}$  triggered events, whilst retaining TODO% of  ${}^8\text{B } \nu_e$  signal events.

### $(\alpha, n)$ Reactions

The impact of  ${}^{238}\text{U}$ - and  ${}^{232}\text{Th}$ -chain isotopes does not simply end at their direct decays. It is possible for the  $\alpha$ s generated during these decays to undergo their own interactions



**Fig. 6.6:** Schematic of the three dominant modes of  $(\alpha, n)$  interaction, taken from [8].

**Fig. 6.7**

with nuclei in the detector. Within the organic scintillator of SNO+, the dominant interaction of this type is when an  $\alpha$  collides with a  ${}^{13}\text{C}$  nucleus, emitting a neutron:  $\alpha + {}^{13}\text{C} \longrightarrow {}^{16}\text{O} + n$ . This is known as an  $(\alpha, n)$  reaction.

The topology of this reaction in the detector is a delayed coincidence, as shown in Fig. 6.6. For the prompt signal, there is the light emitted from the  $\alpha$  just before, and the  $n$  just after the  $(\alpha, n)$ . The neutron then thermalises and gets captured by another nucleus — usually hydrogen in SNO+ — which creates an excited state that then eventually decays, creating a  $\gamma$  that creates the delayed signal in the detector [].

As can be seen in Fig. 6.7,  $(\alpha, n)$  interactions can lead to events reconstructed at a wide variety of energies, which could be an issue for this analysis. However, because they are delayed coincidence events with a typical decay time of  $\sim 100$  ns, the aforementioned out-of-window Bi–Po tagging algorithm also efficiently tags  $(\alpha, n)$  events. Looking again at Fig. 6.7, simply by using the out-of-window Bi–Po tagger without any further modifications TODO% of events in the ROI are cut.

## External Backgrounds

All materials within the SNO+ detector are radioactive, not just the liquid scintillator cocktail. This includes the acrylic, ropes, external water, and PMTs. These components

**Fig. 6.8**

have had their radiopurity “assayed” (that is, measured) throughout the detector’s lifetime, often back to the construction of the original SNO detector itself. The materials other than the liquid scintillator are known to have far higher background levels, especially in the important  $^{238}\text{U}$ - and  $^{232}\text{Th}$ -chain backgrounds [54]. To distinguish between the inherent backgrounds within the scintillator, and the backgrounds from materials at larger radii, we use the terminology “internal” and “external”, respectively.

Although there are numerous external backgrounds, with a suitably accurate and precise position reconstruction algorithm they can be suitably handled. The simplest approach is with a so-called “fiducial volume” (FV) cut: just throw out all events that reconstruct beyond some radius. The only external background events that will reach within the FV are those that have reconstructed very poorly, or have some long-distance radiation that manages to deposit radiation close to the centre of the AV. Because  $\alpha$  and  $\beta$  radiation can only travel short distances through the detector, it is only  $\gamma$  radiation that can realistically travel far enough into the detector to be able to reconstruct anywhere near the centre. Moreover, the intensity of this  $\gamma$  radiation attenuates exponentially towards the centre of the detector, meaning only a tiny fraction of the total number of external events reconstruct within a 3.5 m FV, say. This strong radial-dependence can be seen in Fig. 6.8.

What this figure also demonstrates is that our solar signal has a completely different radial dependence to these backgrounds. As a result, if one considers not just the energy of events but also their reconstructed radius, then it is possible to get an additional handle on the external backgrounds. The FV cut can then be pushed further out to larger radii, allowing one to gain more signal statistics.

Work by Tereza Kroupova [47] allows for additional means of distinguishing external backgrounds from the solar signal. The underlying assumption in the reconstruction of

**Fig. 6.9**

<sup>1</sup> SNO+ events is that there was an electron at a single point, which is entirely valid for  
<sup>2</sup>  $^8\text{B}$  elastic scattering events. However, external backgrounds can fail this assumption in  
<sup>3</sup> two ways. Firstly, these radioactive decays often generate  $\gamma$  radiation in addition to the  
<sup>4</sup> main  $\alpha/\beta$  particle, which creates a multi-site event. Because the `scintFit` position  
<sup>5</sup> reconstruction algorithm is not prepared for a distribution of energy depositions in the  
<sup>6</sup> scintillator, the  $t_{res}$  distribution will broaden. This allows an event classifier to be built  
<sup>7</sup> that distinguishes between the  $t_{res}$  distributions of single-site events and externals,  
<sup>8</sup> known as the “external background timing classifier”. Secondly, because external  
<sup>9</sup> backgrounds that do reconstruct close to the centre of the detector typically have a  $\gamma$   
<sup>10</sup> that travelled a long distance towards the centre of the detector, we expect the earliest  
<sup>11</sup> light that hits the PMTs to arrive most often along the direction of the reconstructed  
<sup>12</sup> position vector. A distribution of PMT hits for a given event as a function of their  
<sup>13</sup> angular distribution relative to the direction of position reconstruction can be built,  
<sup>14</sup> and compared to the expected distributions for single-site and external background  
<sup>15</sup> events. This is known as the “external background topological classifier”. Much like  
<sup>16</sup> the classifier described in Section 6.1.2, the single-site events used for comparison were  
<sup>17</sup>  $0\nu\beta\beta$  events, but these have a similar single-site structure to the solar signal of interest  
<sup>18</sup> in this analysis.

<sup>19</sup> Fig. 6.9 shows the correlation between the two classifier results for both a typical  
<sup>20</sup> external background, and  $^8\text{B} \nu_e$  events.

<sup>21</sup> **Cosmogenic Isotopes**

<sup>22</sup> The final source of background events are radioactive isotopes that form via collisions of  
<sup>23</sup> cosmic rays with atomic nuclei, known as cosmogenic isotopes. Most of these isotopes  
<sup>24</sup> are short-lived [], with lifetimes  $\mathcal{O}(1\text{s})$ . Fortunately, the depth of SNO+ means that

---

6.1 Analysis Methodology

83

our rate of cosmic ray muons interacting with the detector is only 3 an hour [1]. Because the rate is so low relative to other experiments, relatively straightforward approaches to tagging and removing cosmic ray muons and their cosmogenic followers can be utilised without substantial loss of livetime. Events are tagged as a cosmic ray muon if they create a sufficient number of hits in the outward-looking PMTs above background levels, as well as many hits within the detector itself. The details of this tagging for the scintillator-fill were put in place by Lorna Nolan [2], modifying the existing algorithms used in the water-phase [3] and in SNO [4]. After a tagged cosmic muon event, all events for the next 20 s are then vetoed as a means of rejecting followers. This simple cut is enough to remove the vast majority of expected cosmogenic events in the scintillator-phase. The expected impact on loss of livetime, and hence quantity of signal events, is 3 lots of 20 s vetoes an hour, that is to say 1/60<sup>th</sup> of the signal is lost via this cut.

There is one cosmogenic isotope with a long-enough half-life that even the 20 s muon follower veto is not sufficient to remove all events. This is <sup>11</sup>C, which  $\beta^+$ -decays to stable <sup>11</sup>B with a half-life of 20.4 min. The maximum possible energy deposited in this decay is 1.982 MeV [5], just below  $E_{\min}$ , so only a small fraction of <sup>11</sup>C events end up in the ROI: the ones with very high energies that get their energy reconstruction falsely-inflated by some amount. As a result, this background is expected to be very much sub-dominant to all other backgrounds in this analysis. Because this background is important to consider for some other analyses, a triple-coincidence tagging algorithm is being built by Katharine Dixon [6], but not used for this analysis currently.

---

**6.1.3 The Log-likelihood Test Statistic**

At the highest level, this analysis involves taking the data observed in the scintillator-fill after applying a certain set of cuts, along with simulated PDFs for all processes believed

<sup>1</sup> to build up the observed data with those same cuts applied, and then attempting to fit  
<sup>2</sup> the combined energy and radial distributions of the MC to that of the data. Given a  
<sup>3</sup> set of PDFs, to try and match the distribution of observables in data we can modify  
<sup>4</sup> a number of parameters. These consist of the normalisations of each PDF (i.e. the  
<sup>5</sup> total number of events observed due to that process), and any systematic parameters  
<sup>6</sup> that could modify the shapes of these distributions. For this analysis, the neutrino  
<sup>7</sup> oscillation parameters act as *de facto* systematic parameters, as they modify the shape  
<sup>8</sup> of the  ${}^8\text{B}$  PDFs. Of course, unlike usual systematics the oscillation parameters are  
<sup>9</sup> what we are actively trying to measure, instead of being a nuisance.

<sup>10</sup> In order to perform a fit to data in this way, we must first answer a set of questions:

- <sup>11</sup> 1. Which signal and background processes must we consider?
- <sup>12</sup> 2. In addition to their normalisations, are there any further parameters necessary  
<sup>13</sup> to specify the distributions of the PDFs for each of the processes? Systematics  
<sup>14</sup> and oscillation parameters are good examples.
- <sup>15</sup> 3. What is our test statistic?
- <sup>16</sup> 4. What algorithm do we use to try and find the best-fit result?
- <sup>17</sup> 5. How do we measure uncertainties on these best-fit values for each parameter?

<sup>18</sup> In section 6.1.2, question 1 was answered for this analysis. We now give the answer to  
<sup>19</sup> question 3; all other questions on this list will be answered shortly.

<sup>20</sup> The test statistic used for this analysis is the *binned extended log-likelihood*. Once  
<sup>21</sup> the data and MC PDFs have been binned in both the observables of interest, it is  
<sup>22</sup> assumed that the expected number of events in a given bin  $j$  is governed by a Poisson  
<sup>23</sup> distribution:

$$\text{P}_j(n_j|\lambda_j) = \frac{\lambda_j^{n_j} e^{-\lambda_j}}{n_j!}, \quad (6.1)$$

where  $P_j(n_j|\lambda_j)$  is the probability of observing  $n_j$  events in bin  $j$ , given an expectation of  $\lambda_j$  events in total from signal and background processes in that bin. This  $\lambda_j$  can be decomposed into each of the expected rates for each process,  $i$ :

$$\lambda_j = \sum_{i=1}^{N_{\text{PDFs}}} \mathcal{N}_i P_{ij}(\boldsymbol{\theta}), \quad (6.2)$$

where  $N_{\text{PDF}}$  is the number of PDFs being considered in the analysis,  $\mathcal{N}_i$  is the normalisation parameter of the  $i^{\text{th}}$  PDF, and  $P_{ij}(\boldsymbol{\theta})$  is the probability of observing an event of process type  $j$  in bin  $i$ , assuming a set of non-normalisation parameters  $\boldsymbol{\theta}$ . By combining the probabilities of all the bins together, the total probability for a given set of processes assuming these parameters to give rise to the data seen is:

$$P(\mathbf{n}|\mathcal{N}, \boldsymbol{\theta}) = L(\mathcal{N}, \boldsymbol{\theta}|\mathbf{n}) = \prod_{j=1}^{N_{\text{bins}}} \frac{\left[ \sum_{i=1}^{N_{\text{PDFs}}} \mathcal{N}_i P_{ij}(\boldsymbol{\theta}) \right]^{n_j} e^{-\sum_{i=1}^{N_{\text{PDFs}}} \mathcal{N}_i P_{ij}(\boldsymbol{\theta})}}{n_j!}, \quad (6.3)$$

where  $N_{\text{bins}}$  is the total number of bins being considered in the analysis. This probability can be re-framed as the likelihood of the vectors of parameters  $\mathcal{N}$  and  $\boldsymbol{\theta}$  given the vector of number of events in each bin,  $\mathbf{n}$ :  $L(\mathcal{N}, \boldsymbol{\theta}|\mathbf{n})$ . It is rare to see the likelihood as-is, instead, for computational purposes the log-likelihood is used instead,  $\mathcal{L}(\mathcal{N}, \boldsymbol{\theta}|\mathbf{n}) := \ln L(\mathcal{N}, \boldsymbol{\theta}|\mathbf{n})$ . We can then get to the formula actually used for this analysis:

$$\mathcal{L}(\mathcal{N}, \boldsymbol{\theta}|\mathbf{n}) = - \sum_{i=1}^{N_{\text{PDFs}}} \mathcal{N}_i + \sum_{j=1}^{N_{\text{bins}}} n_j \ln \left( \sum_{i=1}^{N_{\text{PDFs}}} \mathcal{N}_i P_{ij}(\boldsymbol{\theta}) \right). \quad (6.4)$$

#### 6.1.4 The Bayesian Statistical Approach

There are two main schools of statistical inference, “Frequentist” and “Bayesian”. In the former, probabilities describe the fraction of times a situation can be found within the whole ensemble of possible worlds. For the latter, we care not about an ensemble

<sup>1</sup> of worlds but instead our degree of belief in this current one. We update our beliefs as  
<sup>2</sup> we acquire knowledge of the world through Bayes' Theorem:

$$\text{<sup>3</sup>} P(\boldsymbol{\mu}|\mathbf{x}) = \frac{\mathcal{L}(\boldsymbol{\mu}|\mathbf{x}) P(\boldsymbol{\mu})}{P(\mathbf{x})}. \quad (6.5)$$

<sup>4</sup> Here,  $\boldsymbol{\mu}$  is the set of parameters that model our system,  $P(\boldsymbol{\mu})$  is our *prior* (pre-existing)  
<sup>5</sup> distribution for those model parameters, and  $\mathbf{x}$  is the data taken in our experiment. The  
<sup>6</sup> updated, *posterior* distribution  $P(\boldsymbol{\mu}|\mathbf{x})$  is then the prior multiplied by the likelihood  
<sup>7</sup> of parameters  $\boldsymbol{\mu}$  given observations  $\mathbf{x}$ ,  $\mathcal{L}(\boldsymbol{\mu}|\mathbf{x})$ , and divided by the total probability  
<sup>8</sup>  $P(\mathbf{x})$  of observing  $\mathbf{x}$  under any circumstance.

<sup>9</sup> Both approaches to statistics are widely-used in statistical analysis, in both particle  
<sup>10</sup> physics and beyond. The Bayesian approach was used for this analysis, as it was  
<sup>11</sup> believed that this helps keep transparent what assumptions are being made in the  
<sup>12</sup> analysis.

<sup>13</sup> Now, if one is able to determine the overall posterior distribution, then it is possible  
<sup>14</sup> to derive best-fit values with uncertainties for all parameters in the fit. This is done by  
<sup>15</sup> “marginalising” the posterior distribution, i.e. integrating over all parameters other  
<sup>16</sup> than the one of interest. A sensible best-fit value is then the modal marginalised  
<sup>17</sup> posterior density, the highest value in the marginalised distribution. The uncertainty  
<sup>18</sup> on this value is derived from the spread of the marginalised posterior, by the calculation  
<sup>19</sup> of the  $1\sigma$  Credible Interval (CI): this is the set of values for a given parameter which  
<sup>20</sup> has a total posterior probability of 68.3%, and contain the best-fit value. There are in  
<sup>21</sup> fact an infinite number of CIs that satisfy this property; for this analysis, the values  
<sup>22</sup> are chosen in decreasing order of marginalised posterior probability density.

### 6.1.5 Markov Chain Monte Carlo

Of course, all of this assumes that one can accurately determine the posterior density distribution. Whilst the likelihood and prior distribution are straightforward enough to calculate, often-times  $P(\mathbf{x})$  (which acts as a normalisation) is very challenging to determine. This is because calculating this normalisation involves integrating the likelihood over all the parameter space, and if there are a large number of parameters this can become enormously numerically complex. An alternative approach is needed!

That alternative comes in the form of *Markov Chain Monte Carlo*, MCMC. A Markov Chain is any mathematical system for which the next state of the system is dependent only on its current state; the system is in some sense “memoryless”. For a large class of Markov Chains — those that are “ergodic” and “aperiodic” — one can prove that regardless of the initial position on the chain, the probability distribution converges to the same distribution  $\pi$ . MCMC uses such a Markov Chain which attempts to converge towards the posterior density distribution in particular. In MCMC, after choosing the initial position in the parameter space, successive states are chosen at random with a probability dependent only on the properties of the current position in parameter space and the proposed position. The convergence property of Markov Chains means that the set of steps made in the parameter space after some initial number of steps will have a distribution that converges to that of the posterior density distribution.

There are a number of MCMC algorithms, and the particular one used in this analysis is that of the *Random-Walk Metropolis-Hastings Algorithm*. In this algorithm, after the initial position in the parameter space  $\mu$ , a new step is proposed some distance from the current one,  $\mu'$ . This step is chosen at random from a multivariate Gaussian distribution centred on the current position, with widths in each dimension of the parameter space chosen beforehand as constants for tuning the MCMC process. This

<sup>1</sup> choosing of a new proposed step at random is what gives the algorithm its Monte Carlo  
<sup>2</sup> and Random Walk titles. Once a new step is proposed, it is accepted as the new step  
<sup>3</sup> with a probability  $S(\boldsymbol{\mu}'|\boldsymbol{\mu})$  according to the condition of *detailed balance*:

$$\begin{aligned} \text{<sup>4</sup>} \quad S(\boldsymbol{\mu}'|\boldsymbol{\mu}) &= \min \left( 1, \frac{P(\boldsymbol{\mu}'|\mathbf{x})}{P(\boldsymbol{\mu}|\mathbf{x})} \frac{R(\boldsymbol{\mu}|\boldsymbol{\mu}')}{R(\boldsymbol{\mu}'|\boldsymbol{\mu})} \right) = \min \left( 1, \frac{L(\boldsymbol{\mu}'|\mathbf{x}) P(\boldsymbol{\mu}')}{L(\boldsymbol{\mu}|\mathbf{x}) P(\boldsymbol{\mu})} \frac{R(\boldsymbol{\mu}|\boldsymbol{\mu}')}{R(\boldsymbol{\mu}'|\boldsymbol{\mu})} \right) \\ \text{<sup>5</sup>} \quad &= \min \left( 1, \frac{R(\boldsymbol{\mu}|\boldsymbol{\mu}')}{R(\boldsymbol{\mu}'|\boldsymbol{\mu})} \exp \left[ \mathcal{L}(\boldsymbol{\mu}'|\mathbf{x}) - \mathcal{L}(\boldsymbol{\mu}|\mathbf{x}) + \ln \frac{P(\boldsymbol{\mu}')}{P(\boldsymbol{\mu})} \right] \right). \end{aligned} \quad (6.6)$$

<sup>7</sup>  $R(\boldsymbol{\mu}'|\boldsymbol{\mu})$  is the probability density that position  $\boldsymbol{\mu}'$  is proposed as a step from position  $\boldsymbol{\mu}$ ,  
<sup>8</sup> and vice versa for  $R(\boldsymbol{\mu}|\boldsymbol{\mu}')$ . In most cases, because of the use of the same multivariate  
<sup>9</sup> Gaussian in choosing proposals,  $\frac{R(\boldsymbol{\mu}|\boldsymbol{\mu}')}{R(\boldsymbol{\mu}'|\boldsymbol{\mu})} = 1$  simply. This component only becomes  
<sup>10</sup> important at the edges of the parameter space, preventing the sampling probability  
<sup>11</sup> being incorrectly impacted if a proposed step goes outside the allowed parameter space.

<sup>12</sup> It is the detailed balance condition that ensures convergence of the MCMC algorithm  
<sup>13</sup> to specifically the posterior density distribution. Crucially, because it is only dependent  
<sup>14</sup> on the ratio of posterior densities, the hard-to-calculate normalisation  $P(\mathbf{x})$  in both  
<sup>15</sup> posterior density terms cancels out, meaning one only needs to calculate the likelihood  
<sup>16</sup> and priors for each step, as well as  $\frac{R(\boldsymbol{\mu}|\boldsymbol{\mu}')}{R(\boldsymbol{\mu}'|\boldsymbol{\mu})}$ .

<sup>17</sup> The specific implementation of MCMC used for this analysis is that of **OXO**, a  
<sup>18</sup> C++ analysis framework first developed by Jack Dunger [50]. **OXO** is able to run  
<sup>19</sup> the Metropolis-Hastings algorithm on multidimensional binned data, using the log-  
<sup>20</sup> likelihood defined in 6.1.3. This framework also allows one to include systematic  
<sup>21</sup> parameters that can float within the fit, and define non-uniform priors for normalisations  
<sup>22</sup> and systematics that have constraints: the details of this will be discussed shortly.

### 6.1.6 Choosing Priors

For this analysis, the suggestions made by Biller & Oser in [55] about choosing prior distributions are followed: for parameters that do not have some pre-existing constraint, a flat prior is used. A nice consequence of this choice is that  $\ln \frac{P(\mu')}{P(\mu)} = 0$ , so the actual value of the prior for these variables never needs to be calculated when running the MCMC algorithm. For the bulk of this analysis, uniform priors are assumed on the neutrino oscillation parameters  $\Delta m_{21}^2$  and  $\theta_{12}$ , as the magnitudes of these parameters are now well-established.

For parameters with existing asymmetric constraints  $\beta_{-\sigma_-}^{+\sigma_+}$ , this analysis uses an asymmetric Gaussian prior, equivalent to the logarithm of the prior being an asymmetric quadratic:

$$\ln P(\mu) = \mathcal{A} - \begin{cases} \frac{(\mu-\beta)^2}{2\sigma_+^2} & \text{if } \mu \geq \beta, \\ \frac{(\mu-\beta)^2}{2\sigma_-^2} & \text{if } \mu < \beta. \end{cases} \quad (6.7)$$

Here,  $\mathcal{A}$  is the logarithm of the prior's normalisation constant, and cancels out in the detailed balance condition. For parameters with symmetric constraints,  $\sigma_+ = \sigma_-$ , then  $\ln P(\mu)$  reduces to a quadratic with maximum at  $\mu = \beta$ .

### 6.1.7 Including Systematics in the Fit

One important implementation detail is how systematics are applied within the MCMC fitting process. Once systematics are added to the fit, at every step the binned PDFs for all the processes considered in the fit must get modified appropriately, which can become extremely computationally-intensive if not approached carefully. The strategy used in the `OXO` framework starts by thinking of the contents of a binned PDF as a vector of bin probabilities,  $\mathbf{p} = (p_1, p_2, \dots, p_{N_{\text{bins}}})^T$ . Then, we can think of a systematic acting on the PDF as a linear transformation, and hence a matrix  $S$  acting

**Fig. 6.10**

on this vector:  $\mathbf{p}' = S\mathbf{p}$ . We only need to calculate this matrix once for a given set of systematic parameter values, and can then use the same matrix on all the PDFs in the fit. Furthermore, when multiple systematics are applied, the matrix for each systematic can then be combined via matrix multiplication into one single “detector response” matrix. OXO uses the `Armadillo` [] linear algebra package for efficient matrix manipulation.

There is a problem that can arise when considering the impact of systematics near the edge of the analysis ROI. Many systematics such as shifts, scalings, and convolutions use information about the contents of nearby bins to determine the contents of a particular bin. However, for bins near the edge some of that information does not exist — it has been lost to the cuts that define the ROI. This can lead to a bias in the generation of the modified PDFs, and therefore also the posterior distribution.

As an example, consider the impact of an energy scale systematic on the energy distribution of  $^{234m}\text{Pa}$  events in the detector, shown in Fig. 6.10. Because the events seen for this process in the ROI are merely the high-energy tail, any systematic energy scaling  $E'_{\text{reco}} := \beta E_{\text{reco}}$  should have a large impact on the number of events observed at the low end of the ROI. However, given that the information about data below  $E_{\text{min}}$  is lost to the ROI cuts, any energy scaling of  $\beta > 1$  will not be applied correctly at all.

The solution to this problem is defining a “buffer region” of bins on either side of the ROI, which allow for tracking of events in and out of the ROI due to systematics, but aren’t considered when calculating the likelihood. This is also shown in Fig. 6.10. After the scaling systematic is applied, although incorrect bin values are found in the buffer region, this is fine because we are no longer calculating the likelihood with those bins. Note that because of this modification, the normalisation parameters we put into the model no longer represent the expected number of events in the ROI. Instead, they

represent the number of events expected in both the ROI and buffer region, before any systematics have been applied.

### 6.1.8 Including Oscillations in the Fit

Within the analysis MCMC code, the process of neutrino oscillations are thought of as a *de facto* systematic that acts only on the  ${}^8\text{B}$   $\nu_e$  and  $\nu_x$  signal spectra. Three parameters relevant to the signal are floated within the MCMC fit:  $\Delta m_{21}^2$ ,  $\theta_{12}$ , and  $\Phi_{8\text{B}}$ , the unoscillated  ${}^8\text{B}$  flux relative to the expected rate. For the two signal PDFs, a third “bookkeeping” dimension is added on top of reconstructed energy and radius: the true neutrino energy,  $E_\nu$ . This is necessary for correctly applying oscillations, as the oscillation probability is a function of the neutrino’s energy, not the scattered electron’s. Before the fit, these 3D PDFs are given normalisations corresponding to the expectation of the number of events for each type,  $\nu_e$  and  $\nu_x$ , after cuts but before oscillations have been applied. Strictly speaking there should be zero  $\nu_x$  events before neutrino oscillations: the pre-oscillation rate used here is the post-cut number of events expected if 100% of the neutrinos oscillated to the  $\nu_x$  type.

During the MCMC fit, for a given set of parameters  $\boldsymbol{\theta} = (\Delta m_{21}^2, \theta_{12}, \Phi_{8\text{B}})$  the following is performed to oscillate the signal PDFs. Firstly, the normalisations are scaled by the factor  $\Phi_{8\text{B}}$ . Then, for each  $E_\nu$  bin the survival probability  $P_{ee}(E_\nu, \Delta m_{21}^2, \theta_{12})$  is calculated. Each bin then has their probability scaled by either  $P_{ee}$  or  $1 - P_{ee}$ , for  $\nu_e$  and  $\nu_x$  respectively. Of course, within the structure of the **OXO** framework these bin-by-bin scaling aren’t immediately applied, but instead a matrix describing the impact of oscillations on each of the PDFs is made. Because the oscillation transformation is purely a bin-by-bin scaling, the resulting matrices are diagonal, with diagonal elements  $\Phi_{8\text{B}} \cdot P_{ee}(E_\nu, \Delta m_{21}^2, \theta_{12})$  or  $\Phi_{8\text{B}} \cdot (1 - P_{ee}(E_\nu, \Delta m_{21}^2, \theta_{12}))$  for  $\nu_e$  and  $\nu_x$  respectively. After the oscillation matrix along with all other systematic matrices are applied to the

<sup>1</sup> signal PDFs, the PDFs are then marginalised over the  $E_\nu$  dimension so that the signal  
<sup>2</sup> PDFs match the dimensionality of all other processes.

<sup>3</sup> Calculations of the survival probability are handled with **PSelmaa**, an algorithm  
<sup>4</sup> written by Nuno Barros for the SNO 3-phase Analysis []. This considers not only the  
<sup>5</sup> neutrino oscillations through the vacuum of space between the Sun and Earth, but also  
<sup>6</sup> the impact of matter effects in both the Sun and Earth. This can usually be a very  
<sup>7</sup> computationally-intensive process, but **PSelmaa** takes advantage of the assumption  
<sup>8</sup> that the solar oscillation parameters are in the so-called “Large Mixing Angle” regime,  
<sup>9</sup> making the calculation much faster. As seen in Section 1.2.1, previous solar oscillation  
<sup>10</sup> experiments demonstrate that this assumption is reasonable. For this analysis, the  
<sup>11</sup> standard MSW effect is assumed with neutrinos obeying the Normal Hierarchy, with  
<sup>12</sup> the Sun following the **B16\_GS98** metallicity model [] and the **PREM** model being used  
<sup>13</sup> for the Earth [].

<sup>14</sup> One final thing **PSelmaa** needs to know to calculate survival probabilities is the  
<sup>15</sup> distribution of solar zenith angles during the data-taking. The solar zenith  $\theta_z$  is the  
<sup>16</sup> angle between the two following vectors: one going from the centre of the Earth through  
<sup>17</sup> the centre of the SNO+ detector, and another starting from the detector’s centre and  
<sup>18</sup> pointing towards the Sun. As an example, if the Sun were ever to be directly overhead  
<sup>19</sup> the detector, both vectors would be along direction  $\hat{z}$  in detector coordinates, leading  
<sup>20</sup> to a solar zenith angle of  $\theta_z = 0$ . The position of the SNO+ detector on Earth, as  
<sup>21</sup> well as the times at which the detector was live, determine the solar zenith angle  
<sup>22</sup> distribution. If not accounted for, this can lead to a bias in the result of the analysis,  
<sup>23</sup> as a preponderance of livetime taken at night (say) would lead to a larger fraction of  
<sup>24</sup> solar neutrinos having to pass through the bulk of the Earth to get to the detector,  
<sup>25</sup> and hence the impact of the MSW effect would be greater.

Even after using the Large Mixing Angle approximation, having to call `PSelmaa` numerous times for every step in the MCMC algorithm would lead to exorbitant run times for the fitting. Therefore, a further approximation is made. Before running the MCMC fit, `PSelmaa` is used to calculate  $P_{ee}$  over the necessary 3D space of parameters. To get a fine scan of this space, 101  $E_\nu$  values from 1 MeV to 20 MeV, 101  $\Delta m_{21}^2$  values from  $3 \times 10^{-6}$  eV<sup>2</sup> to  $1 \times 10^{-3}$  eV<sup>2</sup>, and 151 values for  $\theta_{12}$  from 5° to 65° were looked over. This 3D grid of  $101 \cdot 101 \cdot 151$   $P_{ee}$  values is then written to disk, and loaded into memory for use during the fit as a lookup table. At run-time, as the Metropolis-Hastings algorithm samples this 3D space the survival probability is estimated through a trilinear interpolation of the 3D grid loaded in: a version of linear interpolation for three dimensions [].

## 6.2 Analysis on Scintillator-Phase data

### 6.2.1 Dataset and Livetime

The data used in this analysis was chosen to be scintillator phase data after the end of the PPO top-up campaign that completed in April 2022. Following an initial validation of the analysis tools was performed on data between 29<sup>th</sup> April and 10<sup>th</sup> May 2022 [], data for the full analysis was chosen from runs taken after these dates. Not all data taken during this time was considered usable for this analysis, however: this is the role of *run selection*. The ‘Preliminary Scintillator Gold’ run selection list was used as the basis for this analysis. This run list requires:

- The run type must be in ‘Physics’, as opposed to running in calibration or maintenance;
- The run must last at least 30 minutes;

- 1     • Detector electronics must be working in a stable manner without any alarms,
- 2         and with all crates online;
- 3     • There are no abnormal rates of tagged muons, and the OWL PMTs are correctly
- 4         functioning;
- 5     • There are no unusual conditions from e.g. earthquakes, blasting activity in the
- 6         mine, or loss of power.

7     In the end, data for this analysis used ‘Gold’ runs selected between 17<sup>th</sup> May–30<sup>th</sup>  
8     November 2022, run numbers 300733–306498. The raw livetime associated with this  
9     dataset was calculated by looking at the start and end times of each run using the  
10    detector’s 10 MHz clock: 84.977 days.

11    Importantly, the livetime of the data actually used in the analysis ends up being  
12    somewhat less than this raw value. This is because two of the cuts used in the analysis  
13    (described in Section 6.2.2), the muon tagger and the high-nhit event tagger, also veto  
14    events for a time period following a tagged event. This means that any events in such a  
15    time window are automatically thrown out of consideration for analysis. An algorithm  
16    was written to determine the loss of livetime from both of these tagging processes for  
17    all the runs selected in this analysis, allowing for the net livetime to be calculated. To  
18    ensure accuracy in the value of this lost livetime, the algorithm took care to handle any  
19    overlaps in the time veto windows. This ends up being a quite common occurrence, as  
20    tagged muon events and their followers often have a very high nhit. The net livetime  
21    was calculated to be 80.615 days. Data processing and simulations used RAT versions  
22    7.0.8 and 7.0.9.

## 6.2 Analysis on Scintillator-Phase data

Cut Description	Cut
Event triggers detector	<code>evIndex</code> $\geq 0$
Pass data cleaning	<code>ANALYSIS_MASK</code> = 0x2100000042c2
High-nhit veto	<code>correctedNhits</code> $\geq 5000$ ; veto for 20 s
ScintFitter used	<code>scintFit</code> = true
Results of ScintFitter valid	<code>fitValid</code> = true
In energy Region of Interest	$2.5 \text{ MeV} < E < 14.0 \text{ MeV}$
Fiducial Volume cut	$R < 5.0 \text{ m}$
Remove BiPo Out-of-window tags	See Section 6.1.2
Remove BiPo In-window tags	Tagged if $\text{alphaBeta212} > 0$ and $\text{alphaBeta214} > 0$
Remove External tags	Tagged if $E < 5.0 \text{ MeV}$ , <code>ext0NuTimeTl208AVNaive</code> $> -0.007$ , and <code>ext0NuAngleTl208AV</code> $> -4.7$

**Table 6.1:** Cuts used in this analysis.

### 6.2.2 Event Selection

Once the runs of processed data and matching MC have been selected, the next step was to perform cuts on both in order to obtain our analysis ROI. The full list of cuts used are shown in Table 6.1, most of which have already been explained in Section 6.1.2.

Most of the other cuts used in this analysis that haven't yet been discussed come from handling differences between data and MC. In MC, it is possible to simulate an event but not have it trigger the detector. In this case, the stored index of that event (`evIndex`) will be less than zero. We never see these ‘events’ in data of course, but they do need to be removed from simulation. On the flip side, there are types of events that we observe in the detector but do not simulate. Many of these events are caught during ‘data cleaning’, which looks for a wide variety of problems in the data when the data is being processed. Any subset of the checks made during data cleaning can be chosen to consider in a given analysis by defining a specific ‘analysis mask’; a number that specifies for each of its binary digits whether to consider the data cleaning process associated with that bit. The mask chosen used in this analysis is the Collaboration’s current standard for the scintillator phase, 0x2100000042c2, which corresponds to:

- **00 cut:** Tags every event which has a GTID ending in 00 in hexadecimal, which is associated with a long-standing issue in the detector triggering system.
- **Junk cut:** Tags events that have channels with multiple recorded hits for that event. These are associated with so-called ‘orphaned’ hits that the event building process doesn’t know what to do with.
- **OWL cut:** Tags events that have at least three OWL PMTs hit, a sign that the event came from a cosmic-ray muon.
- **Polling cut:** Tags all events that occur during ‘polling’, a detector monitoring process.
- **CAEN cut:** Tags all events that have missing CAEN trace data.

- **Muon cut:** Tags muon-like events.

- **Muon follower cut:** Tags all events coming up to 20 s after a tagged muon event.

Beyond the formal data cleaning cuts, a high-**nhit** veto was used to tag and remove instrumental backgrounds known as ‘flashers’, in which a PMT will generate large amounts of light in the detector. After any event with a sufficiently large number of hits, both that event and any following it for 20 s are removed. This method is likely to be fairly conservative, and there exist different methods for cutting flashers within data cleaning. However, these were originally coordinated for the water phase, and hadn’t been updated for the scintillator phase at the time of writing.

The results of all these cuts can be seen for the data in Table 6.2. 654 out of the 18,112,495,770 events within the dataset pass all cuts.

- **List final set of cuts chosen for analysis, along with any explanations for cuts that haven’t already been motivated earlier (e.g. data cleaning).** These are:

## 6.2 Analysis on Scintillator-Phase data

Cut	# Events Remaining	Cut Efficiency (%)
Before cuts	18,112,495,770	100.00
Event triggers detector	18,112,495,770	100.00
Data Cleaning	17,663,711,108	97.52
high-nhit veto	17,212,760,839	97.45
ScintFitter used	10,540,054,756	61.23
ScintFitter results valid	3,466,208,344	32.89
Energy cut	134,252	0.00387
FV cut	2155	1.61
BiPo OOW cut	831	38.56
BiPo IW cut	723	87.00
Externals cut	654	90.46

**Table 6.2:** Impact of each cut on the quantity of events in data.

MC process	# Physics Events Simulated	# Triggered Events Simulated	# Remaining after c
AV $^{214}\text{Bi}$			

**Table 6.3:** Impact of each cut on the quantity of events in data.

- Event index cut (prevents MC events that don't trigger detector from being used) 1
- Data cleaning cuts 2
- High-nhit event timing veto cut 3
- Valid scintFit reconstruction 4
- Reconstructed energy  $2.5 < E < 14$  5
- Radial fiducial volume cut 6
- BiPo out-of-window tag 7
- BiPo in-window classifier cut 8
- Externals classifier cut 9
- “Cleanliness” cut 10
- Position FOM cut 11
- Position FOM cut 12

- 1     • Show impact of cuts on data and MC. Show tables (the full details maybe in an  
2         appendix) indicating this.
- 3     • Describe the finalised choice of binning for PDFs and data.

4 [4 pages]

### 5 6.2.3 Expected Rates and Constraints

- 6     • Show expected rates calculation for both signal and background.
- 7     • Describe the constraints chosen to apply to the fit, and why they can be justified.

8 These are:

- 9         – B8 flux constraint
- 10         – U-238 and Th-232 constraints from BiPo tagging
- 11         – Alpha-n constraint from Po-210 tagging by Serena and Shengzhao
- 12         – External constraints from Tony Zummoo's water-phase externals analysis.
- 13     • Describe the systematics to be added to the fit (just energy scale for now, maybe  
14         also energy smearing?). For other possible systematics, such as those in position,  
15         my aim is to explain why they are sub-dominant and so don't need to be added  
16         to the fit. Will cover more about impact of systematics in the next subsection.

17 [4 pages]

### 18 6.2.4 Results

#### 19 Fit Validation

- 20     • Show plots of parameter values versus step, to demonstrate that the step sizes  
21         have been tuned sufficiently.

---

6.2 Analysis on Scintillator-Phase data

99

- Show auto-correlation plots, to motivate a sensible “burn-in” size. 1
- Posterior density plots for each nuisance parameter, to check that they all look sensible and have sufficient statistics. 2
- Show plot of correlation coefficients between parameters, and note any strongly-correlated parameters. 3
- Show plot of correlation coefficients between parameters, and note any strongly-correlated parameters. 4
- [6 pages] 5

**Oscillation Fit Results**

- Look at the data versus ‘best-fit’ MC plot in energy, radius, and both. (Recall that in MCMC, the ‘best-fit’ is not the point of parameter scape reached at the end, but the point of highest posterior density). Is there a good fit to data? Any clear disagreements? 8
- Show 2D contour plot for oscillation parameter posterior density. Note salient features. Show 1D posterior densities for each oscillation parameter. Derive measurement result for  $\theta_{12}$ , i.e. point of highest posterior density, with uncertainties given by the  $1\sigma$  credible interval. 9
- Show 2D contour plot for oscillation parameter posterior density. Note salient features. Show 1D posterior densities for each oscillation parameter. Derive measurement result for  $\theta_{12}$ , i.e. point of highest posterior density, with uncertainties given by the  $1\sigma$  credible interval. 10
- Show 2D contour plot for oscillation parameter posterior density. Note salient features. Show 1D posterior densities for each oscillation parameter. Derive measurement result for  $\theta_{12}$ , i.e. point of highest posterior density, with uncertainties given by the  $1\sigma$  credible interval. 11
- Show 2D contour plot for oscillation parameter posterior density. Note salient features. Show 1D posterior densities for each oscillation parameter. Derive measurement result for  $\theta_{12}$ , i.e. point of highest posterior density, with uncertainties given by the  $1\sigma$  credible interval. 12
- Show 2D contour plot for oscillation parameter posterior density. Note salient features. Show 1D posterior densities for each oscillation parameter. Derive measurement result for  $\theta_{12}$ , i.e. point of highest posterior density, with uncertainties given by the  $1\sigma$  credible interval. 13
- Show 2D contour plot for oscillation parameter posterior density. Note salient features. Show 1D posterior densities for each oscillation parameter. Derive measurement result for  $\theta_{12}$ , i.e. point of highest posterior density, with uncertainties given by the  $1\sigma$  credible interval. 14
- Show 2D contour plot for oscillation parameter posterior density. Note salient features. Show 1D posterior densities for each oscillation parameter. Derive measurement result for  $\theta_{12}$ , i.e. point of highest posterior density, with uncertainties given by the  $1\sigma$  credible interval. 15

[5 pages]

16

**Impact of Systematics**

- Show impact of modifying certain constraints on the final results of the measurement of  $\theta_{12}$ . In particular: fiducial volume,  ${}^8\text{B}$  flux constraint. 18
- Discussion of systematics post-fit — Hopefully energy scaling parameter should be close to 1, given the Collaboration’s calibration of the optics (light yield in particular). Perform a scan over energy smearing, and check impact of possible radial scale systematic. 19
- Discussion of systematics post-fit — Hopefully energy scaling parameter should be close to 1, given the Collaboration’s calibration of the optics (light yield in particular). Perform a scan over energy smearing, and check impact of possible radial scale systematic. 20
- Discussion of systematics post-fit — Hopefully energy scaling parameter should be close to 1, given the Collaboration’s calibration of the optics (light yield in particular). Perform a scan over energy smearing, and check impact of possible radial scale systematic. 21
- Discussion of systematics post-fit — Hopefully energy scaling parameter should be close to 1, given the Collaboration’s calibration of the optics (light yield in particular). Perform a scan over energy smearing, and check impact of possible radial scale systematic. 22
- Discussion of systematics post-fit — Hopefully energy scaling parameter should be close to 1, given the Collaboration’s calibration of the optics (light yield in particular). Perform a scan over energy smearing, and check impact of possible radial scale systematic. 23

<sup>1</sup> [8 pages]

## <sup>2</sup> 6.3 Sensitivity Projections

- <sup>3</sup> • Using the same production MC, generate PDFs with the expected normalisations
- <sup>4</sup> for longer periods of livetime: 1, 3, and 5 years. This still assumes a scintillator-fill
- <sup>5</sup> with identical detector conditions on average, so not considering the impact of
- <sup>6</sup> BisMSB loading at any point.
- <sup>7</sup> • Describe any further constraints to be assumed on top of the existing analysis
- <sup>8</sup> of data: expected improved constraints on various backgrounds, as well as a
- <sup>9</sup> possible energy-scale calibration constraint from the AmBe source or an internally-
- <sup>10</sup> deployed source, for example.
- <sup>11</sup> • Run MCMC fits to these Asimov PDFs for each livetime scenario. Describe
- <sup>12</sup> results in terms of the improvement to the sensitivity to  $\theta_{12}$  as a function of
- <sup>13</sup> livetime.
- <sup>14</sup> • Also consider scenario of lower backgrounds! How does that impact the result?
- <sup>15</sup> • Could consider scenario of BisMSB deployment, leading to substantially greater
- <sup>16</sup> light yield and hence energy resolution. Given time it would take for Production
- <sup>17</sup> MC to be generated for this scenario, I would likely have to come up with
- <sup>18</sup> something clever and quick to actually do this. There quite possibly won't be
- <sup>19</sup> enough time.

<sup>20</sup> [8 pages]

<sup>21</sup> [61 PAGES TOTAL]

# Chapter 7

## Conclusions and Suggestions for Future Work

- Say what has been achieved in this thesis! In particular:
  - Substantial improvement to the SMELLIE generator in terms of speed and dynamic range
  - A much stronger understanding of the discrepancies between data and MC in SMELLIE
  - The creation of two analyses of SMELLIE data, designed explicitly around being robust to these systematics
  - A measurement of the extinction length of scintillator *in-situ* with SMELLIE at 375 nm, monitored over time
  - A first measurement of the scattering length of the scintillator *in-situ*, monitored over time
  - The creation of an analysis of  ${}^8\text{B}$  solar neutrinos in the scintillator phase to measure the solar neutrino oscillation parameters

- <sup>1</sup> – The first measurement of  $\theta_{12}$  using  $^8\text{B}$  neutrinos in SNO+
- <sup>2</sup> – Projections of this solar analysis' precision at longer livetimes
- <sup>3</sup> • Give suggestions for further work that could be done on both SMELLIE and the
- <sup>4</sup> solar oscillation analysis:
  - <sup>5</sup> – Inclusion of LABPPO's polarisability anisotropy in the detector's optical
  - <sup>6</sup> scattering simulation, and determination of its impact on both Physics and
  - <sup>7</sup> SMELLIE measurements
  - <sup>8</sup> – Further investigation of SMELLIE's wavelength-dependence of the beam
  - <sup>9</sup> profiles, looking both at possible origins for the phenomenon and correcting
  - <sup>10</sup> for this in simulation
  - <sup>11</sup> – Various computational improvements to allow for faster MCMC run-times,
  - <sup>12</sup> including the calculation of the systematic matrices, and possible use of
  - <sup>13</sup> GPUs to parallelise some parts of the computation.
  - <sup>14</sup> – Inclusion of additional solar neutrino components at lower energy into the
  - <sup>15</sup> fit, e.g. Be7. Maybe the addition of lower energies also helps to naturally
  - <sup>16</sup> constrain backgrounds within the fit?
  - <sup>17</sup> – Looking at the impact of various advanced background-rejection procedures,
  - <sup>18</sup> such as event directionality or topology. How much do they help with the
  - <sup>19</sup> sensitivity?
  - <sup>20</sup> – Looking at the impact of splitting data into day and night parts, to try
  - <sup>21</sup> and provide further constraints on any matter effects. Not expected to be
  - <sup>22</sup> significant, so was ignored for this analysis so far.
  - <sup>23</sup> – Performing a combined fit with the reactor anti-neutrino oscillation analysis,
  - <sup>24</sup> which allows for the handling of correlated uncertainties, such as detector
  - <sup>25</sup> response systematics.

[3 PAGES TOTAL]

1



# References

- [1] **SNO Collaboration**, J. Boger *et al.*, The Sudbury Neutrino Observatory, Nuclear Instruments and Methods in Physics Research Section A **449**, 172 (2000). 1  
2  
3
- [2] **SNO+ Collaboration**, V. Albanese *et al.*, The SNO+ experiment, Journal of Instrumentation **16**, P08059 (2021). 4  
5
- [3] **SNO+ Collaboration**, M. Anderson *et al.*, Optical calibration of the SNO+ detector in the water phase with deployed sources, Journal of Instrumentation **16**, P10021 (2021). 6  
7  
8
- [4] **SNO+ Collaboration**, M. Anderson *et al.*, Measurement of the B 8 solar neutrino flux in SNO+ with very low backgrounds, Physical Review D **99**, 012012 (2019). 9  
10  
11
- [5] **SNO+ Collaboration**, A. Allega *et al.*, Evidence of Antineutrinos from Distant Reactors using Pure Water at SNO+, Physical Review Letters **130**, 091801 (2023), arxiv:2210.14154. 12  
13  
14
- [6] **SNO+ Collaboration**, M. Anderson *et al.*, Search for invisible modes of nucleon decay in water with the SNO+ detector, Physical Review D **99**, 032008 (2019). 15  
16
- [7] **SNO+ Collaboration**, A. Allega *et al.*, Improved search for invisible modes of nucleon decay in water with the SNO+ detector, Physical Review D **105**, 112012 (2022), arxiv:2205.06400. 17  
18  
19
- [8] I. Morton-Blake, *First Measurement of Reactor Antineutrinos in Scintillator at SNO+ and Study of Alternative Designs for Large-Scale Liquid Scintillator Detectors*, DPhil Thesis, University of Oxford, 2021. 20  
21  
22
- [9] J. Paton, *First Observation of Directionality in the SNO+ Scintillator Phase*, DPhil Thesis, in preparation, University of Oxford. 23  
24
- [10] **SNO+ Collaboration**, A. Allega *et al.*, First Observation of Directionality in a high Light Yield Scintillator Detector, in preparation. 25  
26
- [11] D. J. Auty *et al.*, A Method to Load Tellurium in Liquid Scintillator for the Study of Neutrinoless Double Beta Decay, Nuclear Instruments and Methods in Physics Research Section A **1051**, 168204 (2023), arxiv:2212.12444. 27  
28  
29
- [12] I. Frank and Ig. Tamm, Coherent visible radiation of fast electrons passing through matter, Compt. Rend. Acad. Sci. URSS **14**, 109 (1937). 30  
31

- [13] J. B. Birks, Chapter 3: The Scintillation Process in Organic Materials I, in *The Theory and Practice of Scintillation Counting*, pp. 39–67, Pergamon Press, , second ed., 1967.
- [14] A. LaTorre, New Measurements for the Timing Spectrum of Electrons and Alphas in LAB+PPO, 2016, SNO+ Internal Document DocDB #4081.
- [15] R. Hunt-Stokes, Emission Timing Tuning - 2.2 g / L Optical Model, 2022, SNO+ Internal Document DocDB #7621.
- [16] T. Kaptanoglu, L. Segui, A. Mastbaum, N. Barros, and I. T. Coulter, Optics Overview and Proposed Changes to RAT, 2016, SNO+ Internal Document DocDB #3461.
- [17] T. Kaptanoglu, Documentation for Attenuation Studies, 2022, SNO+ Internal Document DocDB #7608.
- [18] J. B. Birks, Chapter 6: The Scintillation Process in Organic Materials II, in *The Theory and Practice of Scintillation Counting*, pp. 185–234, Pergamon Press, , second ed., 1967.
- [19] S. Riccetto, RAT optics 2.2g/L - after CM\_Aug2022, 2022, SNO+ Internal Document DocDB #7571.
- [20] G. Mie, Contributions to the Optics of Turbid Media, particularly of Colloidal Metal Solutions, *Annalen der Physik* **25**, 377 (1908).
- [21] L. Lorenz, Sur la lumière réfléchie et réfractée par une sphère (surface) transparente, *Oeuvres scientifiques de L. Lorenz* , 403 (1898).
- [22] H. C. van de Hulst, *Light Scattering by Small Particles*Structure of Matter Series (Dover Publications, New York, 1981).
- [23] L. Rayleigh, On the transmission of light through an atmosphere containing small particles in suspension, and on the origin of the blue of the sky, *The London, Edinburgh, and Dublin Philosophical Magazine and Journal of Science* **47**, 375 (1899).
- [24] J. D. Jackson, Section 10.1: Scattering at Long Wavelengths, in *Classical Electrodynamics*, Wiley, , 3rd ed., 1998.
- [25] X. Zhou *et al.*, Rayleigh scattering of linear alkylbenzene in large liquid scintillator detectors, *Review of Scientific Instruments* **86**, 073310 (2015), arxiv:1504.00987.
- [26] N. G. Jerlov and E. S. Nielsen, *Optical Aspects of Oceanography : [Collection of the Papers Presented at the Symposium on Optical Aspects of Oceanography ... Held at the Institute of Physical Oceanography in Copenhagen, 19th-23rd June 1972]* (Academic Press / Symposium on Optical Aspects of Oceanography (1972 : Copenhagen), London ; New York, 1974).
- [27] A. Einstein, Theory of Opalescence of Homogeneous Fluids and Liquid Mixtures near the Critical State, *Annalen der Physik* **338**, 1275 (1910).

- 
- [28] M. S. Smoluchowski, Molecular kinetic theory of gases in critical opalescence state, as well as some related phenomena, *Annalen der Physik* , 205 (1908). 1  
2
- [29] J. Cabannes, Relationship between the degree of Polarisation and the intensity of light Scattered by anisotropic molecules. New determination of the Avogadro constant, *Journal de Physique et Le Radium* **1**, 129 (1920). 3  
4  
5
- [30] J. F. Eykman, Recherches réfractométriques (suite), *Recueil des Travaux Chimiques des Pays-Bas* **14**, 185 (1895). 6  
7
- [31] K. Majumdar, *On the Measurement of Optical Scattering and Studies of Background Rejection in the SNO+ Detector*, DPhil Thesis, University of Oxford, 2015. 8  
9  
10
- [32] E. Turner, *A Measurement of Scattering Characteristics of the Detection Medium in the SNO+ Detector*, DPhil Thesis, University of Oxford, 2022. 11  
12
- [33] H. W. C. Tseung and N. Tolich, Ellipsometric measurements of the refractive indices of linear alkylbenzene and EJ-301 scintillators from 210 to 1000 nm, *Physica Scripta* **84**, 035701 (2011), arxiv:1105.2101. 13  
14  
15
- [34] B. Moffat, *The Optical Calibration of the Sudbury Neutrino Observatory*, PhD thesis, Queen's University, Kingston, Canada, 2001. 16  
17
- [35] E. Hecht, Section 4.6.2: The Fresnel Equations, in *Optics*, Pearson, , fourth ed., 2014. 18  
19
- [36] M. E. Moorhead, *Reflectors in Cherenkov Detectors*, DPhil Thesis, University of Oxford, 1992. 20  
21
- [37] S. D. Biller, N. A. Jelley, M. D. Thorman, N. P. Fox, and T. H. Ward, Measurements of photomultiplier single photon counting efficiency for the Sudbury Neutrino Observatory, (1999). 22  
23  
24
- [38] J. Dunger and E. Turner, Occupancy Analysis for Multiple-Hit Corrections in Large Scintillation Detectors, 2016, NuPhys 2016 poster; SNO+ Internal Document DocDB #4107. 25  
26  
27
- [39] M. Howe *et al.*, Sudbury neutrino observatory neutral current detector acquisition software overview, *IEEE Transactions on Nuclear Science* **51**, 878 (2004). 28  
29
- [40] J. C. Loach, *Measurement of the Flux of 8B Solar Neutrinos at the Sudbury Neutrino Observatory*, DPhil Thesis, University of Oxford, 2008. 30  
31
- [41] M. Dragowsky *et al.*, The 16N calibration source for the Sudbury Neutrino Observatory, *Nuclear Instruments and Methods in Physics Research Section A: Accelerators, Spectrometers, Detectors and Associated Equipment* **481**, 284 (2002). 32  
33  
34  
35
- [42] P. G. Jones, *Background Rejection for the Neutrinoless Double Beta Decay Experiment SNO+*, DPhil Thesis, University of Oxford, 2011. 36  
37

- 1 [43] I. T. Coulter, *Modelling and Reconstruction of Events in SNO+ Related to Future*  
2 *Searches for Lepton and Baryon Number Violation*, DPhil Thesis, University of  
3 Oxford, 2013.
- 4 [44] W. Parker, MultiPDF Method for Position Reconstruction, 2021, SNO+ Internal  
5 Document DocDB #7257.
- 6 [45] W. Parker, Recoordinating ScintEffectiveSpeed & MultiPDF with Updated 2.2  
7 g/l Optics, 2023, SNO+ Internal Document DocDB #7636.
- 8 [46] M. Mottram, Updated functional form energy reconstruction: EnergyRThetaFunc-  
9 tional, 2015, SNO+ Internal Document DocDB #3488.
- 10 [47] T. Kroupová, *Improving the Sensitivity to Neutrinoless Double Beta Decay in*  
11 *SNO+*, DPhil Thesis, University of Oxford, 2020.
- 12 [48] S. S. Wilks, The Large-Sample Distribution of the Likelihood Ratio for Testing  
13 Composite Hypotheses, *The Annals of Mathematical Statistics* **9**, 60 (1938).
- 14 [49] J. Caravaca, SNO+ Sensitivity to Standard Solar Neutrino Oscillations, 2020.
- 15 [50] J. Dunger, *Topological and Time Based Event Classification for Neutrinoless*  
16 *Double Beta Decay in Liquid Scintillator*, DPhil Thesis, University of Oxford,  
17 2018.
- 18 [51] W. T. Winter, S. J. Freedman, K. E. Rehm, and J. P. Schiffer, The B8 Neutrino  
19 Spectrum, *Physical Review C* **73**, 025503 (2006), arxiv:nucl-ex/0406019.
- 20 [52] I. Esteban, M. Gonzalez-Garcia, M. Maltoni, T. Schwetz, and A. Zhou, The fate  
21 of hints: Updated global analysis of three-flavor neutrino oscillations, *Journal of*  
22 *High Energy Physics* **2020**, 178 (2020).
- 23 [53] J. N. Bahcall, M. Kamionkowski, and A. Sirlin, Solar neutrinos: Radiative  
24 corrections in neutrino-electron scattering experiments, *Physical Review D* **51**,  
25 6146 (1995).
- 26 [54] **SNO+ Collaboration**, S. Andringa *et al.*, Current Status and Future Prospects  
27 of the SNO+ Experiment, *Advances in High Energy Physics* **2016**, 1 (2016).
- 28 [55] S. D. Biller and S. M. Oser, Another look at confidence intervals: Proposal for a  
29 more relevant and transparent approach, *Nuclear Instruments and Methods in*  
30 *Physics Research Section A: Accelerators, Spectrometers, Detectors and Associated*  
31 *Equipment* **774**, 103 (2015).