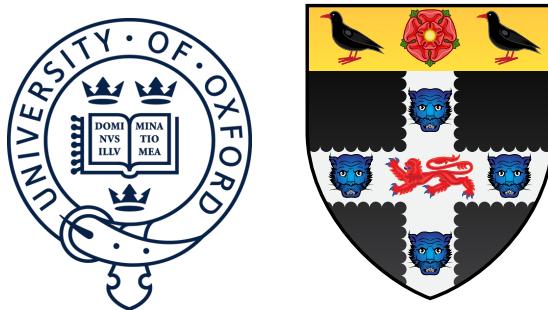


First Measurement of the Solar Neutrino Oscillation Parameters via Boron-8 Solar Neutrinos in SNO+



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A thesis submitted in fulfilment of the requirements for the degree of

Doctor of Philosophy

Trinity Term 2022

To my parents

and

To Oscar Jacobsson:

The best of us

Acknowledgements

And I would like to acknowledge ...

Abstract

Formal 1-page summary of the work completed in the thesis.

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List of Acronyms

AV Acrylic vessel

BD 1,2-Butanediol

BHT Butylated hydroxytoluene

BisMSB 1,4-Bis(2-methylstyryl)benzene

CTC Crate Trigger Card

DAQ Data acquisition (system)

DB Daughter Board

DDA N,N-Dimethyldodecylamine

ECA Electronic Calibration

ELLIE Embedded LED/Laser Light Injection Entity

EXTA External Asynchronous (Trigger)

FEC Front-End Card

FWHM Full-width at half-maximum (of a distribution)

GTID Global Trigger Identification number

GT Global Trigger

LAB Linear alkylbenzene

MTC/A+ Analogue Master Trigger Card

MTC/D Digital Master Trigger Card

npe Number of photoelectrons

OWLs Outward-looking PMTs

PCA PMT Calibration

PMTIC PMT Interface Card

PMT Photomultiplier Tube

PPO 2,5-Diphenyloxazole

PSUP PMT support structure

QE Quantum efficiency (of a PMT)

QHL Charge with high gain over a ‘long’ integration time (390 ns)

QHS Charge with high gain over a ‘short’ integration time (60 ns)

QLX Charge with low gain over a ‘long’ integration time (390 ns)

SNO Sudbury Neutrino Observatory

TAC Time-to-amplitude Converter

TeA Telluric acid, Te(OH)₆

TELLIE Timing subsystem for ELLIE

TeLS Tellurium-loaded liquid scintillator

TTS Transit time spread (of a PMT)

TUBii Trigger Utility Board Mark ii

UPW Ultra-pure water

ZDAB Zebra Database (file format)

RATDB RAT Database

ADC Analogue-to-Digital Converter

Introduction

Couple of pages outlining document's structure and contents (this is what each of the chapters is here for). Less formal than the abstract, also explaining what the expected audience of this thesis is: who will find this document useful!

1

2

3

4

Chapter 1

The Theory of Neutrino Physics

Light

Light

The visible reminder of Invisible Light

The Rock

T. S. ELIOT

1.1 The Standard Model and Neutrinos

1.1.1 A Brief Introduction to the Standard Model

Covering how the SM works at the highest level, including:

- Quantum Field Theory and the Lagrangian dynamical framework
- The connection between symmetries of a QFT model and its gauge fields that describe the model's forces
- The SM's fundamental symmetries, and associated forces, but —

- 1 • Not (exactly) what we see “normally”! The electromagnetic and weak forces
2 appear distinct, and the weak gauge bosons have mass. To explain this, we need
3 a further component, the Brout-Englert-Higgs (BEH) Mechanism.

4 [2 pages total]

5 **1.1.2 Neutrinos within the Standard Model**

- 6 • Basic description of where neutrinos fit into SM: 3 kinds of neutral fermion, the
7 counterparts to the charged fermions. Interacts with the weak force only.
- 8 • Summary of the experimental evidence for this picture: mainly, the discovery of
9 electron anti-neutrinos by Cowan and Reines, the muon neutrino by Lederman,
10 Schwartz, and Steinberger, and the tau neutrino by the DONUT Collaboration.
11 Further critical experiments include the first measurement of a neutrino’s helicity
12 by Goldhaber et al. as well as Danby et al.’s demonstration that ν_μ are distinct
13 from ν_e .

- 14 • More detailed description, via Feynman diagrams, of the two fundamental modes
15 of interaction by neutrinos with the weak force: charged- and neutral-current
16 interactions. A brief mention of the quantitative theory that underlies description:
17 Gashow, Salam, and Weinberg’s Electroweak Theory. This explains not only the
18 V–A structure of charged-current interactions, but also predicted accurately the
19 nature of neutral-current interactions. (Given space constraints, I see no reason
20 to go into much of the details of the theory, or the many experimental tests of
21 its structure.)

22 [4 pages]

1.2 Neutrino Oscillations and Neutrino Masses

1

1.2.1 The Evidence for Neutrino Oscillations

2

- Describe status quo ante of massless nature of neutrinos: BEH mechanism as exists cannot allow for neutrinos to have mass as only left-handed neutrinos have been observed.
3
4
5
- Furthermore, strong experimental limits on neutrino masses, from e.g. tritium-decay endpoint measurements by the KATRIN experiment and cosmological inferences from the CMB by the Planck satellite.
6
7
8
- But — then neutrino oscillations are observed over a variety of experiments and contexts. Summarise critical bits of evidence:
9
10
- Electron neutrino disappearance in solar neutrino experiments, including Ray Davis' Homestake experiment, the SAGE/GALLEX experiments, and SNO. For the latter, the comparison of charged-current and neutral-current modes of interaction was clear evidence of neutrino oscillations over other types of process (e.g. neutrino decay).
11
12
13
14
15
- Include in the above a brief description of Bahcall's Standard Solar Model.
16
- Muon neutrino disappearance in atmospheric and long-baseline accelerator neutrino experiments, such as Super-Kamiokande, T2K, and No ν a.
17
18
- A few further observations to note are: reactor electron anti-neutrino disappearance from both KamLAND and Daya Bay; tau neutrino appearance at the OPERA experiment; short-baseline neutrino anomaly within LSND and MiniBooNE (with recent contrary evidence from MicroBooNE).
19
20
21
22

1.2.2 The Phenomenology of Neutrino Oscillations

- Describe the current phenomenological model of 3-flavour neutrino oscillations that can explain all of this evidence: the PMNS mixing matrix.
- Describe also the MSW effect, which is critical for explaining solar neutrino oscillations.
- Show the formula for solar neutrino oscillations, given this MSW effect in both the Sun and Earth. Note the dependence of solar neutrino oscillations on only the “solar” oscillation parameters. This is all particularly useful for the solar analysis chapter.

$$P_{ee} \left(\tan 2\theta_{12}^M, \sin \theta_{13}^M, \Delta m_{21,M}^2 \right) = BLAH \quad (1.1)$$

[3 pages]

1.2.3 The Origins of Neutrino Mass

- Observed neutrino oscillations require at least two neutrino mass states to be non-zero. Given constraints of the current SM, two main ways of adding neutrino masses: a Dirac mass term (i.e. allowing for sterile neutrinos), and a Majorana mass term.
- For latter, briefly describe what a Majorana particle is, and how with the Seesaw Mechanism (just the simple Type 1 described in-text) one can not only get neutrino masses but also explain their lightness relative to the other massive SM particles. Note that there exist more elaborate versions of this theory.
- Furthermore, with reference to the Sakharov conditions, describe qualitatively how the Seesaw Mechanism also allows for possible leptogenesis/baryogenesis in the early Universe, and hence could explain its matter-antimatter asymmetry.

1.2 Neutrino Oscillations and Neutrino Masses

7

- Describe briefly the nuclear physics behind double-beta decay (i.e. why it can happen at all over just normal beta decay), and then how Majorana neutrinos allow for neutrinoless double beta decay, $0\nu\beta\beta$.
1
2
3
- Describe the experimental signature of $0\nu\beta\beta$: a spike of events of observed energy equal to the Q-value of the decay.
4
5
- Note Schecter-Valle Theorem ensures that any observation of $0\nu\beta\beta$ must be the result of neutrinos being Majorana. I.e. the Universe cannot conspire against us and have $0\nu\beta\beta$ without Majorana neutrinos.
6
7
8
- Very briefly note the current status of the search for $0\nu\beta\beta$, describing the main varieties of experimental setup seen, along with a nice canonical example of such an experiment and their best limit. In particular, the Germanium-crystal detectors such as GERDA, Xenon-TPC detectors like EXO-200, and large-scale liquid scintillators such as KamLAND-Zen.
9
10
11
12
13

[3 pages]

14

[CHAPTER TOTAL: 17 pages]

15

Chapter 2

The SNO+ Detector

The light-soaked days are coming.

JOHN GREEN

2.1 Detector Geometry

The SNO+ detector is a large, multi-purpose neutrino detector built in the SNOLAB underground laboratory near Sudbury, Canada. Its main detector structure is taken from the Nobel prize-winning Sudbury Neutrino Observatory (SNO) [], which can be seen in Fig. 2.1. At the heart of the detector lies the main detector medium, which changes depending on the phase of the experiment — more on the specifics of this in Section 2.2. This medium is held within a 12m diameter sphere known as the Acrylic Vessel (AV). The AV floats within a body of ultra-pure water (UPW), beyond which is a stainless steel support structure (PSUP) that holds 9362 inward-facing Photomultiplier Tubes (PMTs). It is these PMTs that detect the light generated from physics events that occur within the detector medium. The AV is kept in place relative to the PSUP through a series of ‘hold-up’ and ‘hold-down’ tensylon ropes. All of these components are suspended within a large cylindrical cavity also filled with UPW. 91

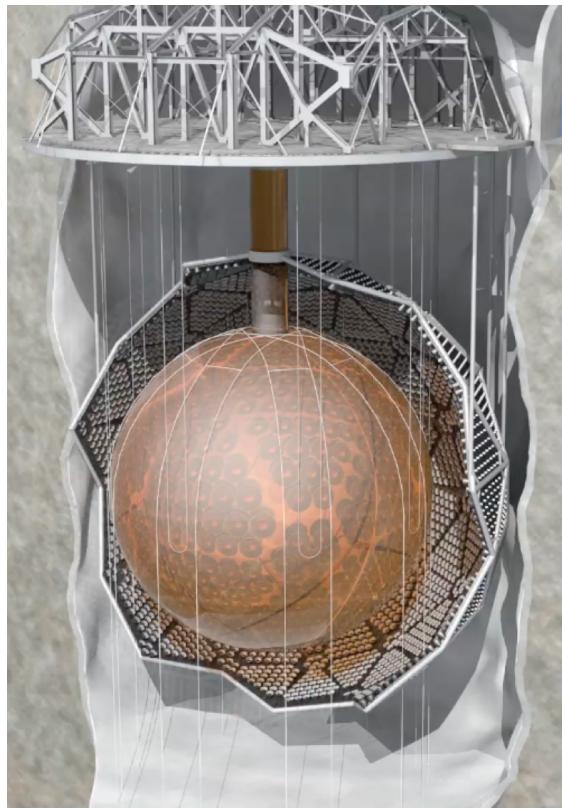


Fig. 2.1: 3D model of the SNO+ detector [1].

1 outward-looking PMTs (OWLs) are also affixed to the outside of the PSUP, allowing
2 for the effective vetoing of cosmic ray muons.

3 Directly above the detector is the Deck, within which all the detector electronics
4 are kept. Access within the AV for calibration tools and filling is possible only through
5 the acrylic ‘neck’ on top of the AV. Full details of the design of the current detector
6 can be found in [1].

7 **2.2 Experimental Phases**

8 As mentioned earlier, SNO+ was designed to fulfil a number of physics goals over
9 multiple ‘phases’ of the detector’s lifetime. The phases are distinguished by the medium
10 that fills the AV. The first main phase (after a brief **Air Fill** phase used only for

2.2 Experimental Phases**11**

detector commissioning) was that of the **Water Fill**, with data taken between May 1
2017 and July 2019. This was used to perform fundamental optical calibrations of 2
the detector [2], measurements of the solar neutrino flux [], observation of neutrino 3
oscillations in reactor anti-neutrinos [], and searches for nucleon decay []. 4

After this, the detector was filled with 800 tonnes of liquid scintillator known as 5
linear alkylbenzene (LAB), mixed with the fluor 2,5-diphenyloxazole (PPO). More 6
information on the physics of scintillators can be found in Section 2.3.1. Filling of the 7
LABPPO cocktail had to be paused in March 2020 due to the COVID-19 pandemic, 8
leading to the detector having its bottom half still filled with UPW, and the top 9
half filled with LAB and PPO at 0.5 g/L. This impromptu phase became known as 10
the **Partial Fill**, and allowed for some creative analyses to be performed: an initial 11
neutrino oscillation analysis from reactor anti-neutrinos [], as well as the first ever 12
observation of directionality in a high light yield scintillator []. Eventually, filling 13
of the detector with liquid scintillator completed in May 2021. At that point, the 14
concentration of PPO in the detector was at 0.6 g/L, markedly below the target level 15
of 2.0 g/L. A further ‘PPO top-up’ campaign then proceeded, finishing in April 2022 16
with a final concentration of 2.2 g/L PPO. Thus began the **Scintillator Fill** of the 17
experiment, which continues on during the time of writing. The main goals for this 18
phase include a number of solar neutrino analyses (including the one described in 19
Chapter 6), a precision measurement of the neutrino oscillation parameter Δm_{21}^2 using 20
reactor anti-neutrinos [3], and further calibrations of the detector and its backgrounds. 21

Finally, in the near future the detector will be loaded with Tellurium for the 22
Tellurium Phase, allowing for the flagship analysis of the experiment to begin: 23
neutrinoless double beta decay. In order to load Te within the liquid scintillator 24
in a stable manner, a chemical loading process has been developed, as described 25
in []. The Te starts within Te(OH)_6 (telluric acid, otherwise known as TeA), which 26

¹ after purification will be reacted with 1,2-butanediol (BD) via heating and addition
² of N,N-Dimethyldodecylamine (DDA), which acts as a stabiliser. What results is
³ tellurium-loaded scintillator, TeLS.

⁴ Two further chemicals are planned to be added to the scintillator cocktail. The
⁵ antioxidant butylated hydroxytoluene (BHT) will be added to capture any free-radicals
⁶ within the liquid scintillator, hopefully preventing any oxidation reactions that could
⁷ lead to the ‘yellowing’ of the scintillator, a degradation of its optical properties. The
⁸ addition of BHT is not expected to impact the detector’s optics in any substantial
⁹ way. However, the other substance to also be added, 1,4-Bis(2-methylstyryl)benzene
¹⁰ (bisMSB), will impact the optics. BisMSB is a ‘wavelength-shifter’ which enables the
¹¹ scintillator cocktail to generate light with a greater overall detection efficiency — more
¹² on the details of this in Section 2.3.1.

¹³ **2.3 Detecting and Recording an Event in SNO+:**
¹⁴ **A Journey**

¹⁵ To understand the SNO+ detector well, it is worth thinking about how the information
¹⁶ contained in a physics event, e.g. a solar neutrino interaction, gets observed. This
¹⁷ section follows the journey of such an event.

¹⁸ **2.3.1 Particle Interactions with Matter**

¹⁹ All observable physics events within the detector begin by the generation of some
²⁰ form of ionising radiation: α , β^\pm , γ , p or n . These can be created via numerous
²¹ processes, both exciting (e.g. $0\nu\beta\beta$ or interactions of neutrinos) and annoying (e.g.
²² decay of background radioisotopes): see Section 6.1.2 for some of them. Regardless of
²³ their origin, these particles begin propagating through the detector, and interacting

2.3 Detecting and Recording an Event in SNO+: A Journey

13

with the detector medium. A number of mechanisms then allow for the generation of
1 optical-wavelength light as a result of these interactions.
2

Cherenkov Light Emission

Whenever a charged particle travels through a dielectric medium at speeds faster than
4 the speed of light in that medium, light is generated from the ‘wake’ of induced dipoles.
5 This is known as **Cherenkov light**, a process much akin to the ‘sonic boom’ that
6 occurs when an object travels at supersonic speeds through a medium. This light
7 emanates outwards in a cone along the direction of the charge’s travel; the angle of the
8 cone θ_γ is purely a function of the speed of the charged particle relative to the speed of
9 light in vacuum, β , and the refractive index of the medium $n(\omega)$ at a given frequency
10 ω : $\cos \theta_\gamma(\omega) = \frac{1}{n(\omega)\beta}$. There is then a minimum speed necessary for Cherenkov light to
11 be generated: $\beta_{\min}(\omega) = 1/n(\omega)$.
12

In addition to the characteristic cone shape of the light, the spectrum of the light
13 generated is also distinctive. Igor Tamm and Ilya Frank determined the expected
14 energy dE emitted per unit length travelled by the charged particle, dx , per unit of
15 frequency $d\omega$, as []:
16

$$\frac{\partial^2 E}{\partial x \partial \omega} = \frac{q^2}{4\pi} \mu(\omega) \omega \left(1 - \frac{1}{\beta^2 n^2(\omega)} \right) = \frac{q^2}{4\pi} \mu(\omega) \omega \left(1 - \left[\frac{\beta_{\min}(\omega)}{\beta(\omega)} \right]^2 \right), \quad (2.1) \quad 17$$

under the condition that $\beta(\omega) > \beta_{\min}(\omega)$. Here, q is the charge of the moving particle
18 and $\mu(\omega)$ is the permeability of the medium.
19

All SNO+ detection media allow Cherenkov light to be generated, as long as
20 sufficiently high energy particles traverse it. In the water fill phase of the detector,
21 Cherenkov light was the only means by which light could be generated. Light from
22 Cherenkov emission can still be created in liquid scintillator, but it tends to be swamped
23 by another form of light generation: scintillation.
24

1 Scintillation

2 For certain special classes of material, the excitation and ionisation of atomic electrons
 3 nearby a moving charged particle can lead to the generation of optical-wavelength
 4 light, in a process known as **scintillation** (often generally referred to as ‘luminescence’
 5 or ‘fluorescence’). Although multiple varieties of scintillator exist, the one used in
 6 SNO+ is that of an organic liquid scintillator. For such liquids, scintillation light is
 7 generated from the de-excitation of delocalised ‘ π -electrons’ that are found within
 8 certain carbon–carbon bonds. A major example of these π -bonds are found in benzene
 9 rings, which are present in LAB, PPO, and bisMSB.

10 Because of this delocalised structure, excited atomic π -electrons can stay in what
 11 is typically the first-excited state for somewhat longer than typical excited states:
 12 lifetimes of $\mathcal{O}(10^{-9}$ s) as opposed to $\mathcal{O}(10^{-12}$ s). This is what gives scintillation light
 13 its characteristic ‘slow’ response relative to the instantaneous light generated by the
 14 Cherenkov process. Moreover, decays from this state can emit light typically in the
 15 optical-wavelength range. In addition to excited electrons, ionised electrons can also
 16 recombine — that is, re-enter atomic orbitals — into various excited states, and then
 17 decay back to the ground state, also allowing for the possibility of scintillation light to
 18 be generated.

19 Because of atomic spin selection rules [], scintillation light typically has, at the very
 20 least, a ‘fast’ and ‘slow’ time component. In SNO+, we currently model emission of
 21 scintillation light from LAB with 4 time components, following the timing distribution
 22 $f(t)$ given by:

$$23 \quad f(t) = \sum_i A_i \left(\frac{e^{-t/\tau_i} - e^{-t/\tau_{\text{rise}}}}{\tau_i - \tau_{\text{rise}}} \right), \quad t > 0. \quad (2.2)$$

24 Here, A_i and τ_i correspond to the fraction of light emitted and decay constant for each
 25 component respectively, and τ_{rise} is a common rise time. The current fitted values for

Component	A_i	τ_i [ns]
1	0.665	7.35
2	0.218	5.45
3	0.083	117.5
4	0.0346	425
Rise	–	0.8

Table 2.1: Current values used to model scintillator emission from high-energy electrons in 2.2 g/L LABPPO.

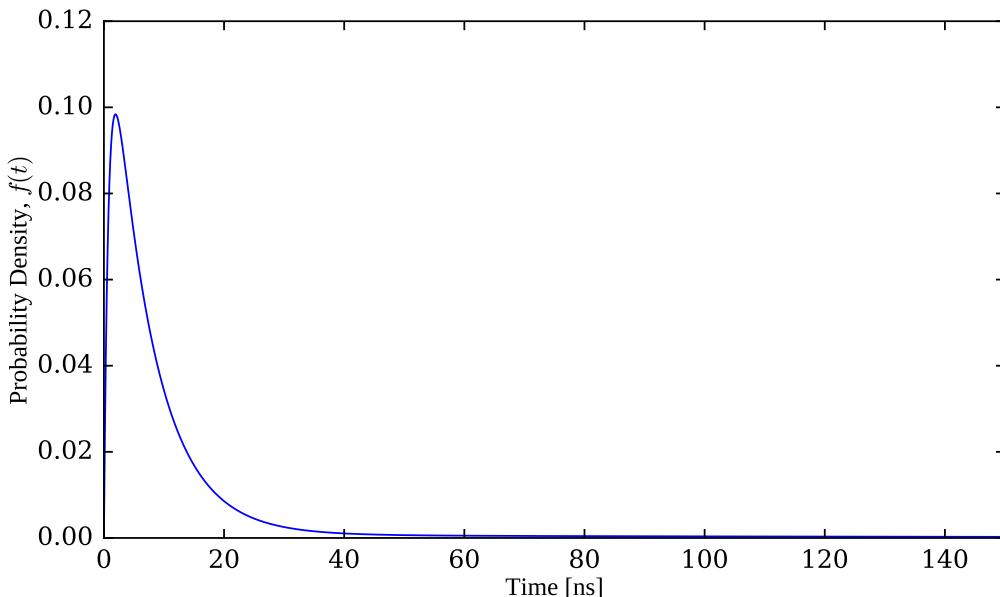


Fig. 2.2: The current modelled distribution for the emission timing of the SNO+ scintillator, for electron tracks [].

these parameters for the emission from electron tracks can be seen in Table 2.1. A plot of $f(t)$ with these values can be seen in Fig. 2.2.

When using just a single scintillating compound, the very same energy levels that can generate scintillation light are those that can absorb it. This can be a problem for large-scale detectors like SNO+, which depend on scintillation light being unobstructed in its path to the PMTs. Conveniently, this problem can be addressed with the addition of another scintillating component, known (somewhat confusingly) as the primary fluor. In SNO+, this is the PPO added to the LAB.

Fig. 2.3

When an LAB molecule is excited, that energy can be transferred to a PPO molecule through what is known as a ‘non-radiative transfer’. In short, this transfer of energy occurs not through the emission and absorption of optical photons, but through the coupling of the molecules’ electric dipoles.¹ When the now-excited PPO molecule de-excites to emit scintillation light, the different molecular structure it has generates a different emission spectrum to that of LAB. These longer wavelengths of light are no longer able to be absorbed by the LAB, allowing for a scintillator with less optical absorption.

Adding in one additional component doesn’t have to be the end, either. In SNO+ we plan on adding in the compound BisMSB to the scintillator cocktail. This is a ‘wavelength-shifter’: scintillation light at short wavelengths is absorbed, and then re-emitted at longer wavelengths, where the detection efficiency of the PMTs is greater. More on the properties of the PMTs in SNO+ can be found in Section 2.3.3. The net effect of the three scintillating components within SNO+ can be seen in Fig. 2.3. Note how, as energy is transferred from one scintillation component to another, the wavelength of light emitted gets necessarily longer as energy is lost to heat.

The light yield of a scintillator, i.e. the amount of optical photons generated per unit of energy deposited into the scintillator, is a function not just of the scintillator but also the incident particle. In particular, α particles are far more effective at exciting and ionising nearby atoms, and so can deposit far more of its energy into the scintillator per unit volume. However, the strength of this ionisation for α s can actually become at detriment to the generation of scintillation light. Empirically, scintillators follow to

¹To be pedantic, photons are still transferred in this energy exchange, but they are virtual instead of real.

2.3 Detecting and Recording an Event in SNO+: A Journey

17

first order Birks' Law for their scintillation light yield []:

$$\frac{dL}{dx} = S \frac{\frac{dE}{dx}}{1 + k_{\text{Birks}} \frac{dE}{dx}}, \quad (2.3)$$

where $\frac{dL}{dx}$ is the number of photons emitted per unit track length, $\frac{dE}{dx}$ is the energy loss of the incident particle per unit track length, S is the scintillator's characteristic light yield constant, and k_{Birks} is the scintillator's "Birks' Constant"². For minimum-ionising particles such as a 6 MeV electron, the denominator of this equation is close to 1, and so the amount of scintillation light generated is just $\frac{dL}{dx} = S \cdot \frac{dE}{dx}$. However, for α -particles generated in radioactive decays, this denominator can become substantial, and in the limiting case we have merely $\frac{dL}{dx} = S$. In the current phase of SNO+, S and k_{Birks} are measured to be $14,000 \gamma/\text{MeV}$ and $0.0798 \text{ mm MeV}^{-1}$,³ respectively [].

2.3.2 Optical Processes

Once optical-wavelength photons have been created within the detector, various processes can then occur that can hinder its path towards a PMT, and therefore modify the observed signal. This subsection covers the main optical processes, with a focus on Rayleigh scattering, as an understanding of this phenomenon is critical for Chapters 3–5.

Rayleigh Scattering

Optical scattering is the general process of how light is scattered by particles within a medium. This is fundamentally an electrodynamical process: an electromagnetic wave is incident on the set of particles within the medium, which induces these particles to

²Birks' Constant is often just written as k_B , but this is easily confused with the far-better known Boltzmann Constant, which is completely different!

³This is the Birks' constant used in simulation for α -particles. Other particles, such as electrons, are given slightly different values of k_{Birks} .

¹ oscillate within the field, and therefore generating their own electromagnetic radiation
² in response. Usually, this ‘scattered’ radiation has the same frequency as that of the
³ incident radiation, and therefore the scattering is said to be *elastic*. It is possible under
⁴ certain circumstances for this scattered radiation to be of a longer wavelength than
⁵ the incident radiation: in which case, energy was absorbed by the particles and so the
⁶ scattering was *inelastic*. However, this latter type of scattering, also known as Raman
⁷ scattering, is not relevant for SNO+ [].

⁸ The general solution to elastic optical scattering was first described by Gustav
⁹ Mie [] and Ludvig Lorenz [] in what is now known as *Mie Theory*. In this theory, it is
¹⁰ assumed that a plane wave of wavelength λ is incident on a dielectric sphere of radius a .
¹¹ While the general solution to the problem of Mie scattering is somewhat complicated
¹² (if tractable), in certain regimes one can make further simplifying assumptions that
¹³ substantially reduce the complexity of the result. In particular, if one assumes that
¹⁴ the size of the particle is much smaller than the wavelength of light, and that any
¹⁵ induced dipole moment can actually be established in the time window allowed by the
¹⁶ oscillation period of the electromagnetic field [], then one can obtain *Rayleigh scattering*.
¹⁷ This simpler case is so-called because of its initial formulation by Lord Rayleigh [].

¹⁸ One can show that the differential cross-section associated with Rayleigh scattering
¹⁹ of unpolarised light off of a single particle, $\frac{d^2\sigma_{\text{Ray}}}{d\theta d\phi}(\theta, \phi)$, is given by []:

$$\frac{d^2\sigma_{\text{Ray}}}{d\theta d\phi}(\theta, \phi) = \frac{8\pi a^6}{\lambda^4} \left(\frac{n_{\text{par}}^2 - 1}{n_{\text{par}}^2 + 2} \right)^2 (1 + \cos^2 \theta). \quad (2.4)$$

²¹ Here, θ and ϕ correspond respectively to the polar and azimuthal angles of the scattered
²² waves, and n_{par} is the refractive index of the scattering particle. Most important to
²³ notice about this equation is that the cross-section follows a strong $1/\lambda^4$ dependence,
²⁴ meaning that short wavelengths of light will be scattered to far greater extents than
²⁵ that of longer wavelengths. Secondly, the light is not scattered isotropically, but

2.3 Detecting and Recording an Event in SNO+: A Journey

19

according to a $1 + \cos^2 \theta$ dependence. This means that most light is either scattered directly forwards or backwards (known as a *backscattering*), and little gets scattered orthogonally to the direction of the incident light. This is useful when it comes to trying to measure scattering in the SNO+ detector, as it provides a handle upon which to distinguish scattered light from isotropically-emitted scintillation light.

Of course, we care about the scattering that occurs within an entire bulk medium, not just the scattering off of a single molecule. From a macroscopic perspective, the key quantity of interest is a material's *Rayleigh scattering length*, l_{Ray} : the mean distance a photon is expected to travel before Rayleigh scattering. One can show that, assuming the above differential scattering cross-section, the Rayleigh scattering length is given by []:

$$l_{\text{Ray}} = \left[\frac{16\pi}{3} R \right]^{-1}. \quad (2.5)$$

R is the *Rayleigh ratio*, $R = \frac{1}{V} \frac{d^2\sigma_{\text{Ray}}(90^\circ)}{d\theta d\phi}$, equivalent to the power of the scattered light per unit volume of the scattering medium per unit incident intensity at $\theta = 90^\circ$.

This can lead to a few changes to Rayleigh scattering that are worth noting. Firstly, unlike for a single particle, the electric polarisability of a material can be *anisotropic*. Anisotropic materials have a modified angular dependence on their differential cross-section, governed by the *depolarisation ratio*, δ . In particular, the $(1 + \cos^2 \theta)$ dependence becomes $(1 + \frac{1-\delta}{1+\delta} \cos^2 \theta)$. For isotropic materials, $\delta = 0$, and so the angular dependence reduces to the original form.

Secondly, the above model has been shown to be insufficient to describe liquids or solids [], because of the non-negligible strength of their inter-molecular forces. Fortunately, Einstein [], Smoluchowski [], and Cabannes [] developed a theory for describing how photons can scatter off of the local charge density fluctuations that naturally are present in a medium because of the thermal motion of molecules. The theory shows that the Rayleigh ratio of a medium is related to the medium's dielectric

¹ constant, ε , by:

$$\text{2} \quad R = \frac{\pi^2}{2\lambda^4} \left[\rho \left(\frac{\partial \varepsilon}{\partial \rho} \right)_T \right]^2 k_B T \kappa_T \frac{6 + 6\delta}{6 - 7\delta}, \quad (2.6)$$

³ where ρ is the density of the medium, $\left(\frac{\partial \varepsilon}{\partial \rho} \right)_T$ is the partial derivative of the dielectric
⁴ constant with respect to a changing density assuming a constant temperature T , k_B
⁵ is the Boltzmann Constant, and κ_T is the medium's isothermal compressibility. This
⁶ latter quantity is given by the rate of change of volume given a changing pressure of
⁷ the medium, all at a constant temperature.

⁸ Furthermore, the Eykman Equation [] has been shown to be an effective empirical
⁹ formula relating how ε is impacted by density fluctuations to the medium's refractive
¹⁰ index, n_{med} :

$$\text{11} \quad \rho \left(\frac{\partial \varepsilon}{\partial \rho} \right)_T = \frac{(n_{\text{med}}^2 - 1)(2n_{\text{med}}^2 + 0.8n_{\text{med}})}{n_{\text{med}}^2 + 0.8n_{\text{med}} + 1}. \quad (2.7)$$

¹² This leads to a final formula for the Rayleigh scattering length:

$$\text{13} \quad l_{\text{Ray}} = \left[\frac{8\pi^3}{3\lambda^4} \left(\frac{(n_{\text{med}}^2 - 1)(2n_{\text{med}}^2 + 0.8n_{\text{med}})}{n_{\text{med}}^2 + 0.8n_{\text{med}} + 1} \right)^2 k_B T \kappa_T \frac{6 + 3\delta}{6 - 7\delta} \right]^{-1}. \quad (2.8)$$

¹⁴ Discussions of the scattering lengths currently assumed within SNO+'s optical
¹⁵ model for UPW and LABPPO can be found within the theses of Krishanu Majumdar [4]
¹⁶ and Esther Turner [5]. In particular, whilst the scattering length of the UPW in the
¹⁷ water phase was measured by Esther, major systematics in the measurement remained.
¹⁸ Measurements of the scattering lengths in scintillator are the focus of Chapters 3–5.

¹⁹ Absorption and Re-emission

²⁰ In addition to scattering, an optical medium is also able to absorb light that propagates
²¹ through it. For a given medium, the *absorption length* l_{abs} is analogous to l_{Ray} described
²² above, and is typically strongly a function of wavelength. For most materials, absorbed
²³ light is forever lost, converted into heat. However, for the special case of scintillators,

2.3 Detecting and Recording an Event in SNO+: A Journey**21**

re-emission of absorbed light is possible: this is because of the physics described in
 1
 Section 2.3.1.
 2

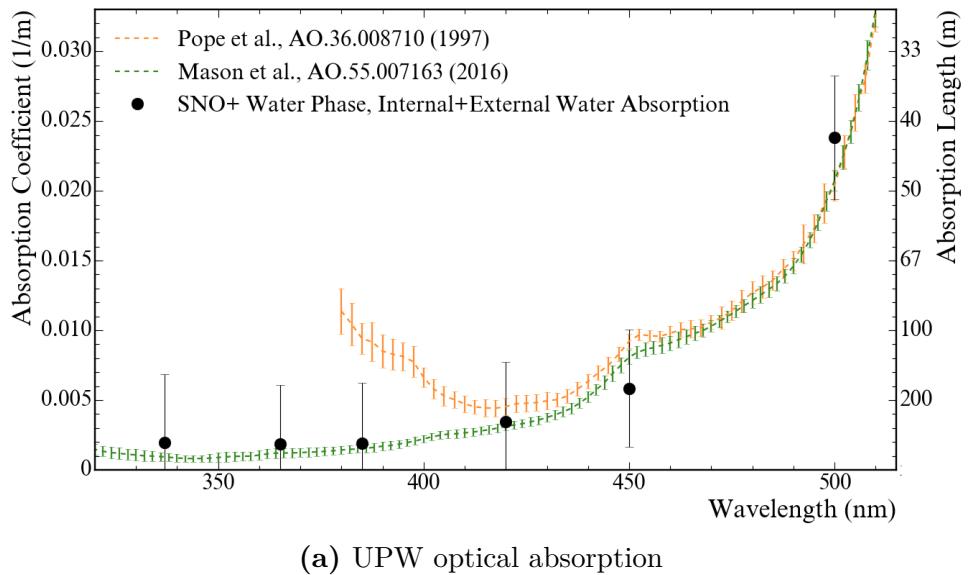
Because both scattering and absorption impede a photon's ability to propagate
 3
 through a medium directly, it is often possible to measure their combined impact
 4
 through what is known as the absorption/extinction length, l_{ext} :
 5

$$\frac{1}{l_{\text{ext}}} = \frac{1}{l_{\text{abs}}} + \frac{1}{l_{\text{Ray}}}. \quad (2.9) \quad 6$$

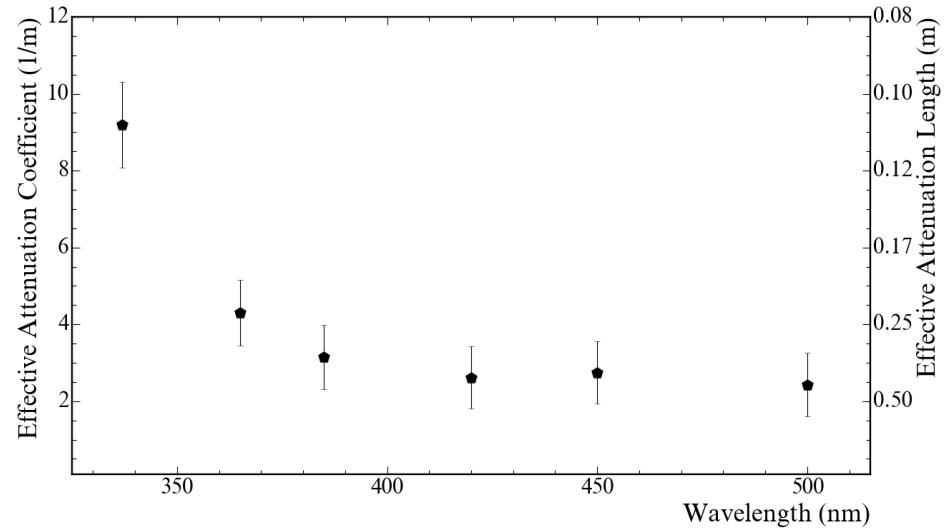
In the water phase, the ‘Laserball’ calibration system was used to measure various
 7
 optical properties of the detector, including the extinction lengths of the UPW and
 8
 acrylic as a function of wavelength [2]. Using the water phase scattering measurements
 9
 made by Esther, Eq. 2.9 allowed for the estimation of the absorption lengths of these
 10
 two materials, shown in Figure 2.4. Measurements of the extinction length in the
 11
 scintillator phase is discussed in detail in Chapter 5.
 12

Surface reflection and refraction

When light travels through the boundary of one medium to another, both reflection and
 14
 refraction can be possible, depending on the relative refractive indices of the two media.
 15
 The refractive indices of the UPW, acrylic, and LABPPO are shown as a function of
 16
 wavelength in Figure 2.5. Note that, for most optical wavelengths, LABPPO has a
 17
 very close refractive index to acrylic, whereas UPW is somewhat farther away. By
 18
 consequence, negligible refraction is expected in most cases for light travelling between
 19
 the liquid scintillator and the acrylic; however, substantial refraction is possible for light
 20
 travelling between acrylic and UPW. Because of this, isotropically-emitting point-like
 21
 physics events within the AV that are close enough to the acrylic will have some of
 22
 their light undergo Total Internal Reflection (TIR) at the AV, reflecting back into the
 23
 AV instead of continuing outward into the outer water.
 24



(a) UPW optical absorption



(b) Acrylic optical attenuation

Fig. 2.4: Measured properties of the UPW and acrylic in the water phase, from [2].**Fig. 2.5**

Even when not undergoing TIR, some light at a boundary can still reflect. The fraction of light that reflects is known as the *reflectance* R , compared to that which is able to transmit through the boundary, the *transmittance* $T = 1 - R$. The *Fresnel Equations* determine the reflectance of an interface [1]:

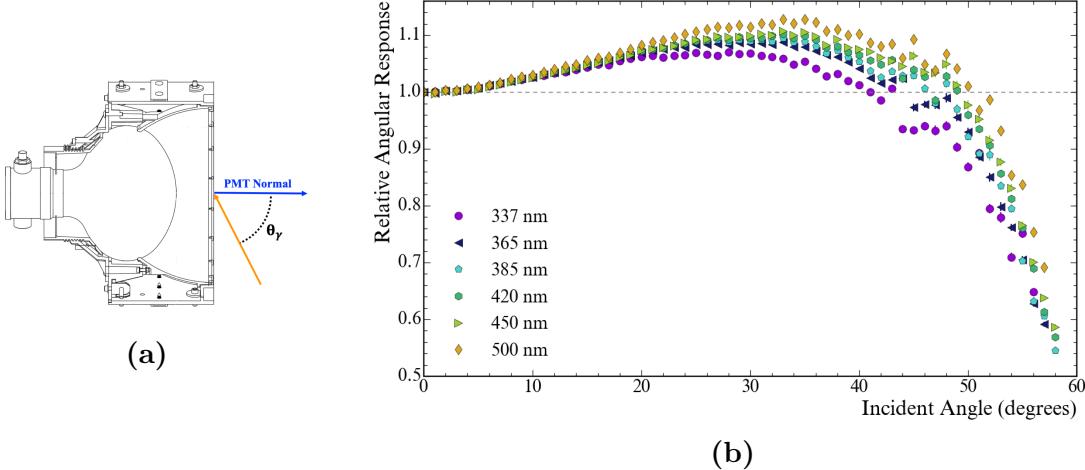
$$R_s = \left| \frac{n_1 \cos \theta_i - n_2 \cos \theta_t}{n_1 \cos \theta_i + n_2 \cos \theta_t} \right|^2, R_p = \left| \frac{n_1 \cos \theta_t - n_2 \cos \theta_i}{n_1 \cos \theta_t + n_2 \cos \theta_i} \right|^2, \quad (2.10)$$

where R_s and R_p are the reflectances of *s*- and *p*-polarised light, n_1 and n_2 are the refractive indices of the first and second optical media, and θ_i and θ_t are the angles of incidence and refraction, respectively. For SNO+, we are only interested in unpolarised light, so the total reflectance $R = (R_s + R_p)/2$.

2.3.3 Detection by PMTs

The final step for photons in our journey is detection by a PMT. Almost all PMTs in SNO+ are of the Hamamatsu r1408 design [2]. These PMTs within SNO+ are housed within a reflecting cone known as a ‘concentrator’. The combined PMT–concentrator ‘bucket’, shown in Fig. 2.6a, is designed to maximise the collection efficiency of light emanating from within the AV, whilst minimising the collection efficiency of light outside the AV [3]. The so-called ‘angular response’ of the PMT buckets has been measured in both SNO and SNO+ using the Laserball, which describes the relative collection efficiency as a function of the polar angle of the incident light ray relative to the direction in which the PMT bucket points. The results of this can be seen in Fig. 2.6b.

Once a photon is incident on the PMT’s photocathode, it is possible for that photon to be absorbed and generate a photoelectron. The probability of this happening is governed by the PMT’s Quantum Efficiency (QE) at the photon’s wavelength, which can be seen in Fig 2.7. Note how the QE is maximal at 420 nm with a value of TODO%.

**Fig. 2.6****Fig. 2.7**

Once this photoelectron has been created, the dynodes within the PMT generate a cascade of electrons that allow for an observable voltage signal to be produced. The dynamics of this cascade are such that there is a natural spread of possible times between the creation of a photoelectron and the generation of the voltage signal pulse in the PMT's wire. This is known as the ‘Transit Time Spread’ (TTS) of the PMTs: for SNO+, the full-width at half-maximum (FWHM) of the TTS for the r1408-type PMTs is TODOns [1].

Finally, if multiple photons generate photoelectrons on the same PMT close enough in time, the detector’s front-end electronics (described in Section 2.3.4) is unable to tell. However, the amount of charge generated increases in proportion to the number of photoelectrons (npe). Much like with the transit time, the strength of the signal observed by the PMT is governed by a distribution, a function of the npe generated. Examples of these distributions can be seen in Fig. 2.8. The relatively large widths of these charge distributions precludes the ability to straightforwardly determine the npe purely from charge when the npe is small. To work around this, various techniques

Fig. 2.8

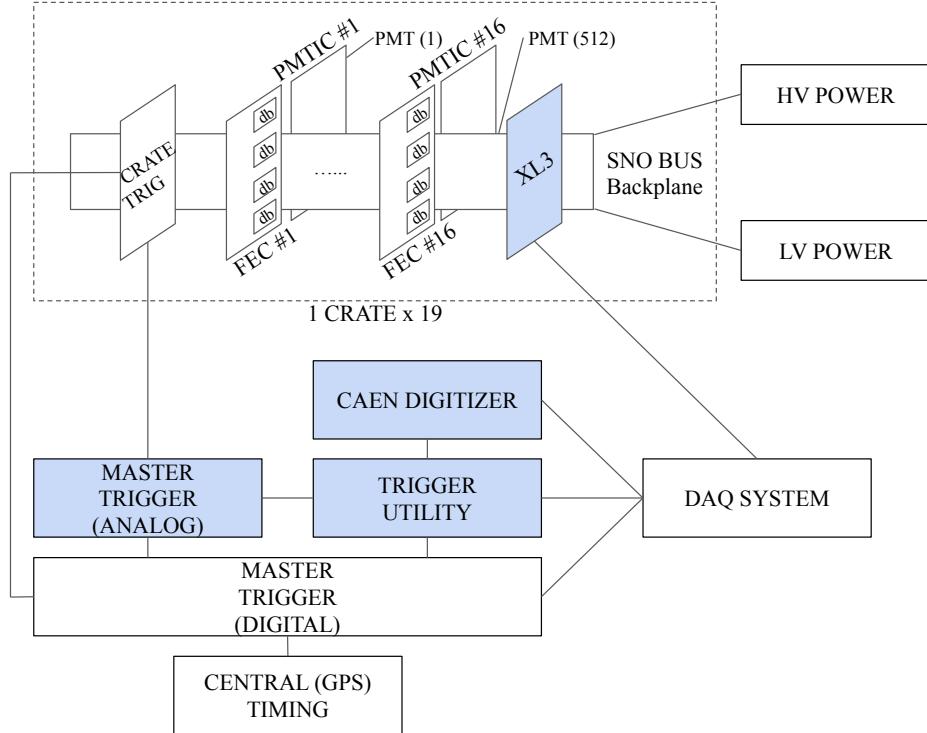


Fig. 2.9: Schematic of the front-end electronics used for data acquisition and triggering in SNO+, taken from and discussed in [1].

can be employed to try and estimate the npe in a given PMT — an example of one such method can be seen in Section 4.2.

2.3.4 Data Acquisition and Triggering

Once a signal is generated in the cable attached to a PMT, it travels along up to the front-end electronics on the deck above the detector. The job of these electronics, known as the data acquisition (DAQ) and triggering system, is to convert raw electronic signals from the PMTs into recorded digital ‘events’ that can be used for analysis. A schematic showing the setup of the electronics is shown in Fig 2.9, with full details in [1].

1 A signal passes first through the PMT Interface Card (PMTIC), which then sends
2 it through to one of the Daughter Boards (DBs) which are stored on Front-End Cards
3 (FECs) within one of 19 electronic crates on deck. The DBs determine if the analogue
4 signal along a given PMT channel has crossed a pre-defined charge threshold, at which
5 point we say that we have detected a ‘hit’ on that PMT’s channel. When this occurs,
6 the DB performs a set of important actions:

- 7 1. Begins a timer for that channel, in the form of a Time-to-Amplitude Converter
8 (TAC). TAC also corresponds to the resulting quantity being measured.
- 9 2. Begins integrating the total charge signal for that channel in three ways, known
10 as QHS, QHL, and QLX. These correspond to using different integration times
11 and gain settings.
- 12 3. Generates trigger pulses for that channel for each available trigger type. Three
13 main trigger signals are the ‘N20’ (a square pulse for 20 ns), ‘N100’ (a square
14 pulse for 100 ns), and ‘ESUMHI’ (a pulse copying the shape of the voltage signal
15 for that channel). The reason for these names will be explained shortly.

16 Whilst the TAC, QHS, QHL, and QLX are being calculated, the trigger signals
17 from each channel are sent over to the Crate’s Trigger Card (CTC), where the signals
18 are then summed for each trigger type. These crate-level trigger signals are then sent
19 over to the 7 detector-level Analogue Master Trigger Card (MTC/A+), which further
20 sum the signals by trigger type from all the crates in the detector. It is at this point of
21 the process where the names of the trigger types becomes clear: the combined N20
22 and N100 signals are proportional to the total number of hit PMTs within a 20 ns and
23 100 ns time window, respectively, whilst the total ESUMHI signal corresponds to the
24 total charge seen over all the PMTs.

If these trigger signals go above certain pre-defined thresholds, then a signal is sent for that trigger type to the Digital Master Trigger Card (MTC/D). The MTC/D receives all trigger signals from the detector, and if a given trigger type has been ‘masked in’ the Card will generate a Global Trigger (GT) for the detector according to its 50 MHz clock. Under certain circumstances, such as calibrations, a trigger signal can be generated externally and asynchronous to the MTC/D: these are ‘EXTA’ triggers. Any such EXTA trigger signal is first handled by an electronics box named ‘TUBii’ (Trigger Utility Board Mark ii), before then being passed onto the MTC/D. This becomes relevant when discussing the calibration electronics described in Chapter 3.

Once a GT signal is generated, it is then sent back to all the CTCs, which then orders the integration of time and charge to be stopped on all channels for that crate. The time and charge information that has been temporarily stored on each channel’s CMOS chip on the FEC is then sent to the crate’s ‘XL3’ card, which packages the crate’s raw information via ethernet over to a set of computers. The trigger signals for a triggered event are also digitised by a CAEN brand Analogue-to-Digital Converter (ADC), and sent to the same DAQ computers. The total window of time in which data is gathered from one GT signal is 400 ns, with data from up to 180 ns before and 220 ns after the GT has arrived. There is then a necessary ‘dead time’ of 420 ns after a given GT has been given in which no further GTs can be made.

Finally, the raw data from the crates and trigger system arrives in a set of computers, which organise all of this into an individually-packaged ‘event’, stored on disk in the ‘Zebra Database’ (ZDAB) format. A given built event contains the TAC and charge information from each hit PMT, the CAEN digitised waveforms, a unique identifying number for that triggered event (the GTID), as well as the times from both the MTC/D’s 50 MHz clock and a GPS-calibrated 10 MHz clock. The former time is used

¹ for measuring relative times between events whilst the latter is used for knowing the
² time of day of an event: both are used in Chapter 6.

³ 2.3.5 Operation of the Detector

⁴ Control of the detector’s DAQ system is handled through a custom-built GUI known
⁵ as ORCA [\[\]](#). This program allows operators of the detector to modify settings in the
⁶ detector electronics at both a low- and high-level. It also allows operators to monitor
⁷ the current status of the detector, such as voltage and current levels within each crate.

⁸ ORCA allows for the detector to have its data split into ‘runs’ of different types.

⁹ For the majority of the time, the detector is run in the ‘Physics’ mode, with individual
¹⁰ runs split into 1 hour periods. It is this data that is used for almost all high-level
¹¹ physics analyses, such as the one described in Chapter 6. Other modes include ones
¹² for detector maintenance, as well as for calibrations of various kinds. During certain
¹³ calibration runs, data can be further split into ‘subruns’ where necessary. Operation of
¹⁴ specific calibration sources, including the one described in Chapter 3, can be performed
¹⁵ through the ORCA GUI.

¹⁶ 2.4 Calibrations and Detector Modelling

¹⁷ Once the raw data from triggered events has been stored in files, certain extra steps
¹⁸ must be taken before effective analysis of that data can be achieved. This section
¹⁹ covers those steps.

²⁰ 2.4.1 Detector Monitoring

²¹ No data taken from the detector can reasonably be used for analysis unless its quality
²² has been approved. This is done in a number of ways on SNO+. Firstly, a number

2.4 Calibrations and Detector Modelling

29

of automated systems monitor all aspects of the detector, including voltage levels
1
in the crates, trigger rates, as well as ‘slower’ quantities such as the tensions on the
2 ropes holding the AV in place. Problems in any of these measured parameters trigger
3 an automatic alarm system, which notifies a human detector operator. A detector
4 operator monitors the detector 24/7 whilst the detector is live.
5

In addition to systems that monitor whether anything has gone wrong, information
6 about the state of the detector during each run is stored in a database known as
7 RATDB. This information includes, amongst other things, a recording of which PMT
8 channels have actually been raised to high voltage for that run, as well as any chan-
9 nels/cards/crates that have been flagged for having a known poor data quality (e.g.
10 being overly noisy).
11

2.4.2 Electronic and PMT Calibrations

The lowest level of calibrations performed in SNO+ are the Electronic and PMT
12 Calibrations: ECAs and PCAs, respectively. These calibrations convert the raw time
13 and charge values recorded by the DAQ into quantities that can actually be used in
14 analysis.
15

During an ECA, two main quantities are measured. Firstly, because of noise the
16 integrated charge measured on each channel is offset by some amount. This offset,
17 known as the ‘pedestal’, is recorded for each channel. The other quantity is the ‘time
18 slope’ for each channel, which allows one to convert from the ADC TAC counts into
19 an uncalibrated hit time of that channel’s PMT in ns. Both of these quantities are
20 measured by sending external signals to channels, forcing them to start measuring TAC
21 and charge even though no PMTs were actually hit. Running ECAs also enable us to
22 spot any channels with unusual behaviour, so that they are not used during analysis.
23 More information about ECAs can be found in [].
24

25

Using ECAs alone is not enough to have fully-calibrated time and charge data. The lengths of cables between PMTs and PMTICs are all slightly different, leading to differences in the so-called ‘cable delay’ of each channel. This means that two PMTs that have a photoelectron generated at the same time can generate slightly different TAC values. Furthermore, because the start time of the TAC is determined by when the channel’s signal goes above a constant threshold, if a signal is very large (e.g. when numerous photoelectrons have been generated on one PMT) then the start time of the TAC will be systematically earlier. This is known as the ‘time walk’. Both of these quantities get measured during PCAs [].

PCAs can be performed by either the Laserball [], or by the TELLIE calibration system []. The latter is a series of 92 optical fibres attached at various points of the PSUP, through which optical-wavelength light can be fired from LEDs. TELLIE is the Timing subsystem of the ELLIE calibration system: the Embedded LED/Laser Light Injection Entity. The other two subsystems, AMELLIE and SMELLIE, are introduced in Section 2.4.3. For both the Laserball and TELLIE calibration systems, the cable delay and time walk are measured by firing light from the source at a known time, and observing when the signal arrives in each PMT channel.

On top of calibrating the PMT hit times, PCAs also further calibrate the charge information. In particular, the charge spectrum generated by a single photoelectron is determined for each channel. This allows us to convert the pedestal-corrected charge ADC counts into a number of photoelectrons.

Using the data gathered from both ECAs and PCAs, the raw data stored in ZDABs is processed into a new file format known as RATDS files. These files contain all the information of an event, but now the times and charge information have been calibrated. It is this file type used in the optical calibration work of Chapters 3–5.

2.4.3 Energy and Optical Calibrations

- After ECAs and PCAs, further calibration is necessary to accurately model the optical properties of the detector. 1
2
3
- Describe briefly the function of the Laserball, SMELLIE, and AMELLIE sources in optical calibration (can reference the water-phase optical calibration paper). 4
5
Details of how each analysis works is obviously not needed here, especially not for SMELLIE as we have 3 chapters to go over that! 6
7
- AmBe and N16 sources provide further information for calibration 8
- In-situ backgrounds used for more calibration sources, especially for determining the light yield of the scintillator, and Birks' constant. 9
10

[2 pages]

11

2.4.4 Event Reconstruction

- Once detector is calibrated, event reconstruction becomes possible. Describe main assumption of a SNO+ event: we assume a single-site electron event. 12
13
14
- Briefly describe the basics of how energy, position, and time reconstruction works, as this is needed for the solar analysis chapter. Mention existence of direction fitting in scintillator! 15
16
17
- ‘Physics’ runs have their events reconstructed in a third pass of data processing, resulting in a ‘fully-processed’ RATDS file as well as a simplified NTUPLE file used in high-level physics analyses (such as my solar analysis). 18
19
20

[2 pages]

21

1 2.4.5 Event Simulation

- 2 • Briefly describe overview of RAT software: not only provides the software by which
3 the above data processing occurs, but allows for a GEANT4-based simulation of
4 events in the detector. This simulation includes all parts of the physics described
5 in Section 2.3, including particle and nuclear physics interactions, optical photon
6 creation and propagation, signal generation and event building based on simulated
7 triggers of the electronics. These simulated events then get processed in the same
8 manner as actual data is.

9 [1 page] [24 PAGES TOTAL]

Chapter 3

The SMELLIE Calibration System

There's a certain Slant of light,

Winter Afternoons —

That oppresses, like the Heft

Of Cathedral Tunes —

EMILY DICKINSON

- Basic principle for how SMELLIE works: firing collimated laser light into detector to observe scattering events.
- Analysis will measure and monitor scattering in a detector with changing optics.
- One can try and measure some component of this: the cross-section/scattering length versus wavelength and time, and/or the relative scattering length versus wavelength and time.

[1 page]

3.1 The SMELLIE Hardware

- Describe the existing hardware, post-upgrade made in Summer 2022. For pre-upgrade hardware, can simply cite previous SMELLIE theses. This includes the path of light into the detector, as well as the path of the trigger signal.
- Make sure to mention explicitly these upgrades: Tony Zummo's fix to the TUBii trigger logic, as well as the addition of the VFA, updated MPU, and modified trigger window. Make sure to motivate why these updates were made.

[7 pages]

3.2 Software for SMELLIE Data-taking

- Can be brief here! Little has changed since previous theses, so can mostly just summarise and cite.
- Server running on SNODROP machine, which converts high-level commands into low-level ones that the hardware can interpret.
- Run plan files written in JSON handed to ORCA which then sends relevant commands to SNODROP which fires as appropriate.
- Operator interacts with ORCA to perform SMELLIE calibration runs.
- After SMELLIE data taken, run description file created, containing metadata about the run conditions, used in analysis.

[2 pages]

3.3 Commissioning SMELLIE in the

Scintillator Phase

- Explain why commissioning of SMELLIE is needed: Need to confirm that SMELLIE is working as expected; determine intensity "set-points" for different use cases.
- Commissioning originally performed by Esther and JeffL back in the water phase; explain why this needed to be re-done for both the scintillator phase and after the hardware upgrades.
- No need to describe the Tesseract in detail here - that can be in Jeff L's thesis. But, I do want to show the results of both commissioning campaigns in scintillator-fill, one before the new hardware was added, and one after.

[5 pages] [15 PAGES TOTAL]

12

Chapter 4

1

Simulating SMELLIE Events

2

Max Power : *Kids. From now on there are three ways of doing things: the right way, the wrong way, and the Max Power way.*

Bart Simpson : *Isn't that just the wrong way?*

Max Power : *Yes, but faster!*

THE SIMPSONS

Critical to extraction of scattering information from SMELLIE data is an accurate Monte Carlo (MC) simulation of the SMELLIE system. By modelling the laser light emission into the detector correctly, we can simulate how SMELLIE light will be impacted by changing scattering lengths in the detector. Because of the complexity of the optics of the optical fibres used to direct the laser light into the detector, a given SMELLIE event is simulated as a partially-collimated “flash” of visible photons emanating from the emission point of the fibre into the detector. This flash then requires a number of parameters to be correctly described. In particular:

4

5

6

7

8

9

10

11

- **Fibre emission positions** were recorded during the installation of the fibres.

12

- **Wavelength and emission timing distributions** of light pulses were taken from measurements of the laser heads by their manufacturers [], or by colleague Jeff Lidgard in the case of the SuperK wavelength distribution [].
- **The “pulse magnitude”**, defined as the mean number of photons simulated per event, is determined on a subrun-by-subrun basis, and is assumed to fluctuate as a Poisson distribution.
- **The beam profiles**, which describe the angular emission distributions of each fibre, is the focus of this chapter. These are necessary because unlike scintillation light, the light emitted from SMELLIE fibres is not isotropic.
- **Nominal fibre emission directions** attempt to define the centre of the beam for a given fibre.

This chapter is split into three sections. Improvements to the existing simulation algorithm for the beam profiles are first made, and then the beam profiles themselves are updated. Finally, comparisons between data and simulation are made after the upgrades to investigate any remaining discrepancies.

4.1 Improving the SMELLIE Generator Algorithm

4.1.1 Previous Attempts at SMELLIE Event Simulation

Before we can determine the beam profiles, we must first decide how to specify them. Previous observations show that different fibres can have notably different beam profiles [4], so we let each fibre’s beam profiles be unique. We assume for now that a given fibre’s beam profile is stable over time, and independent of the wavelength of light fired. A straightforward, naïve approach to parameterising a beam profile would be as follows: specify some nominal fibre direction, corresponding to the direction light

4.1 Improving the SMELLIE Generator Algorithm

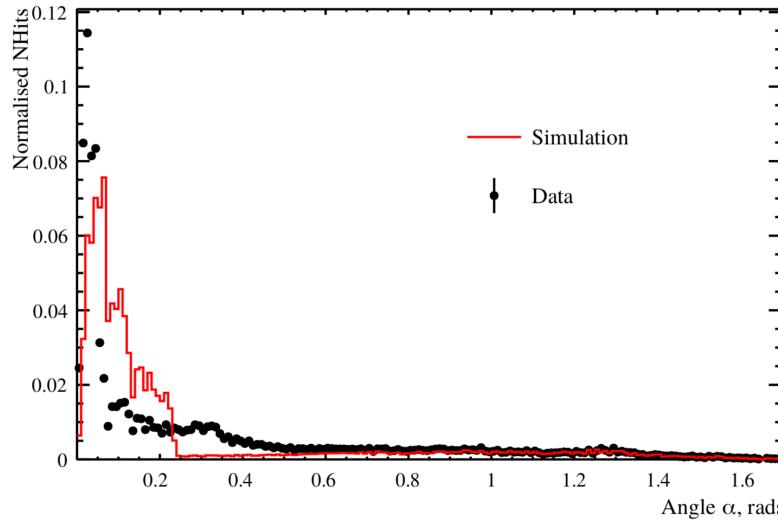


Fig. 4.1: Comparison between a simulation of one of the fibres, made from the 1D beam profile generator (red), with the associated data subrun that was used to create that beam profile (in black). For both MC and data, what is plotted is the PDF of observed PMT hits, as a function of the α angle. Poissonian errors have been added to the data points, but are too small to see. Clearly, this 1D generator does not replicate the observed beam profile correctly. Figure taken from [5].

takes travelling from the fibre to the centre of the “beamspot” observed on the other side of the detector. Then, specify a 1D beam profile, corresponding to the probability density of firing a photon at a given polar angle α relative to the nominal direction. One might even assume this distribution is Gaussian. The distribution in azimuthal direction, ϕ , is assumed to be uniform.

This 1D beam profile approach was used initially for SMELLIE, and remains in use for the other ELLIE sub-systems within SNO+. However, when SMELLIE data was taken in the water-phase of the experiment, simulations using these beam profiles failed to match them well at all - see figure 4.1 for an example. Not only was the distribution in α not Gaussian, a distinct speckle-pattern can be observed within the beamspot that is not uniform in ϕ . This fact led to colleague Esther Turner building a SMELLIE generator that could handle 2D beam profiles: dependent on both α and ϕ . The distribution was stored as a map from each inward-pointing PMT in the detector

¹ to a relative intensity value. This was chosen because the beam profile shapes were
² calibrated from existing SMELLIE data — more on this in section 4.2.

³ This original 2D generator then sampled the beam profile via a rejection sampling
⁴ approach, outlined as follows:

- ⁵ 1. Propose a test direction (α, ϕ) , by generating ϕ uniformly in the interval $[0, 2\pi]$,
⁶ and α according to some pre-determined Gaussian distribution, known as the
⁷ Gaussian envelope.
- ⁸ 2. Given this test direction, calculate where a line following this direction from the
⁹ fibre of interest will hit the PSUP on the other side of the detector. Find the 3
¹⁰ closest PMTs to that point.
- ¹¹ 3. From those PMTs, obtain their relative intensity values from the beam profile
¹² mapping, and perform an interpolation based on how close each PMT is to the
¹³ PSUP intersection point. This gives an interpolated relative intensity value for
¹⁴ this test direction.
- ¹⁵ 4. Because we are sampling using the angular coordinates (α, ϕ) , differential area
¹⁶ elements over this space of directions do not have the same size. We can correct
¹⁷ for this fact by multiplying our interpolated relative intensity by $\sin \alpha$, which
¹⁸ corresponds to the Jacobian of the direction-space.
- ¹⁹ 5. Calculate the value for the Gaussian envelope along this test direction.
- ²⁰ 6. Throw a random number uniformly between 0 and the Gaussian envelope value. If
²¹ the random number is less than the interpolated intensity, then this test direction
²² is accepted, and a photon is generated with that direction. Otherwise, we reject
²³ the direction and try the whole process again.

 4.1 Improving the SMELLIE Generator Algorithm
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This generator certainly works, but has a key problem: efficiency. The 1D generator was able to generate a SMELLIE event (that is, to fully specify the starting parameters of all the photons emitted from a fibre) at a speed of ~ 1 ms. However, the 2D generator specified here could take upwards of ~ 50 s *per event* to generate. Because a typical SMELLIE analysis requires simulating many millions of events, the CPU time taken to perform this quickly became unfeasible. Fixing this generator speed problem was a high priority for the SMELLIE analysis.

4.1.2 The new generator
8

On careful inspection of the existing 2D generator, the main reason for the slowness of the algorithm is the use of a rejection approach. Even with use of the Gaussian envelope, which was included to help with speed, the vast majority of proposed directions are never selected. Figure 4.2 shows a histogram of number of attempts per event it took for a valid direction to be chosen for a representative SMELLIE simulation. Moreover, the calculations needing to be done for every proposed direction are relatively complex, notably trying to find the 3 nearest PMTs to some point on the PSUP.

A new 2D generator was built with these thoughts in mind. Firstly, the rejection method would no longer be used, given its inefficiency. We would also endeavour to try and “pre-calculate” as much as possible before run-time. Starting with the existing PMT relative intensity maps, we plot these in the 2D direction-space ($1 - \cos \alpha, \phi$): see Figure 4.3a. In a toy-MC simulation, 500,000 directions are then thrown uniformly in this 2D space per fibre. For each direction, the same method of obtaining an interpolated intensity value from the nearest PMTs to the corresponding point on the PSUP as from the original 2D generator was performed, the only difference being that these calculations were done well before any actual SMELLIE simulation. Figure 4.3b shows the interpolated intensities obtained for one fibre.

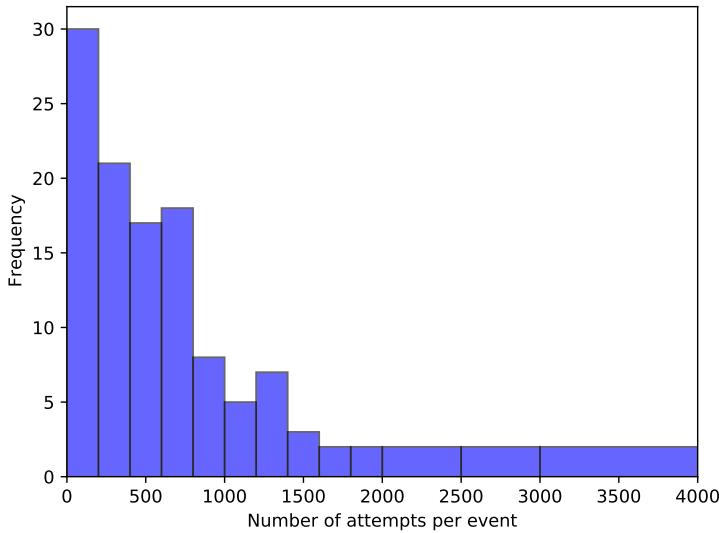


Fig. 4.2: Typical distribution of the number of attempts it takes for the existing 2D generator before the test direction gets accepted, per event.

Following this, the sampled intensities were then binned into a 2D histogram, where the bin value corresponds to the sum of all intensities for all directions found within this bin. Choosing a sensible binning procedure is important: too few bins, and necessary information about the shape of the beam is lost, whilst too many bins can oversample the data and capture statistical artefacts in the sampling process instead of just the beam profile. As a balance, 15 bins were chosen along the ϕ direction, and 60 in $r = 1 - \cos \alpha$. This was chosen to ensure that a reasonable number of PMTs were located within each bin, lessening the impact of any statistical fluctuations. Although the bins in ϕ were chosen to have uniform width, this was decided to be not the case for the other axis, as there is far more important information near $r = 0$ (the beamspot). Instead, the width of the bins in r were calculated so that roughly the same total probability was contained in each r -strip. By consequence, bins near the beamspot typically are of significantly smaller size than ones much further out. This allows us to both capture any rapid changes in intensity near the beamspot, where this

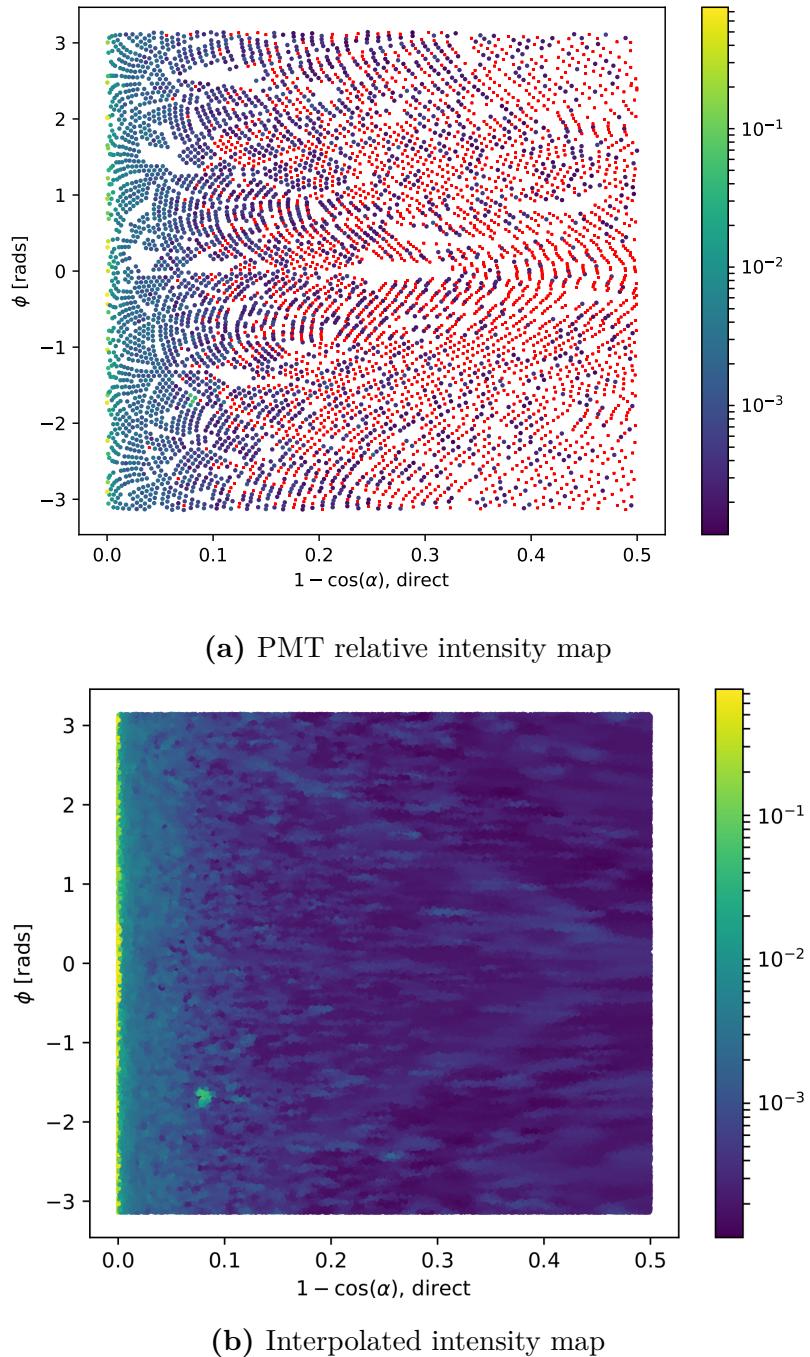


Fig. 4.3: The first step in the new method for preparing the new generator. In (a), the relative intensities used for the existing beam profile of fibre labelled FS055 are shown for each PMT, the position on the plot indicating the location of that PMT in the fibre coordinates. The colour indicates the relative intensity; PMTs marked red have an intensity of zero. Figure (b) shows the result of throwing 500,000 directions uniformly over this 2D space, the intensity of each point given by interpolating the intensities of nearby PMTs.

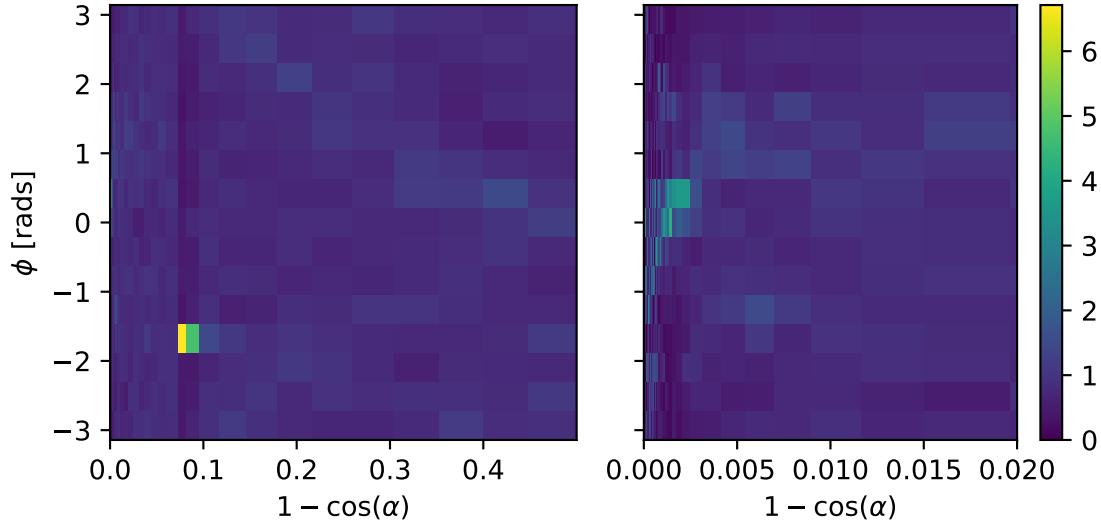


Fig. 4.4: Histogram of interpolated intensities within the 2D direction-space. The left view shows the full histogram; the right is a zoomed-in version near the beamspot. Unlike the binning in ϕ , the bin widths in r are not at all uniform. Instead, they have been determined such that the area summed over a given “strip” of bins of constant r will be the same.

¹ matters greatly, and smooth out the very-low intensities seen at larger polar angles.

² One of these histograms can be seen in Figure 4.4: the large change in bin widths as a
³ function of r is clear. One can also see that near the beamspot notable dependence on
⁴ the intensity as a function of ϕ . The mysterious “spot” at $r = 0.08$, well out of the
⁵ beamspot, is an indication that the underlying beam profile data being used requires
⁶ improvement: more on this in section 4.2.

⁷ The Cumulative Density Function (CDF) of this intensity histogram as a function
⁸ of bin was then produced, where the bins were ordered through a raster-scan: scanning
⁹ first over ϕ , and then r . The CDF was then normalised to 1 so that it was well-defined.

¹⁰ It is this CDF object that is then loaded in and sampled from during event generation.
¹¹ To do this, an “inverse-CDF” approach was used, which has the major benefit over
¹² rejection sampling of always producing a valid direction for every sample made. The
¹³ algorithm works as follows:

- ¹⁴ 1. Throw a random number uniformly in $[0, 1]$.

4.2 Improving the beam profiles

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2. Perform a binary search to find the bin that has the largest CDF value below
this random number. 1
3. Look at the bin edges in ϕ of this selected bin: use linear interpolation of the
random number to obtain a ϕ value located between these two ϕ -values. 2 3
4. Look at the selected bin's r -bin edges, and select a value of r by throwing a
second random number uniformly between the two edges. Convert this r into a
polar angle α . 4 5 6 7
5. The photon's direction is defined by the (α, ϕ) chosen by this process. 8

Because of the relative simplicity of this algorithm compared to the previous 2D generator, the speed improvement was very large: generation now took ~ 1 ms per SMELLIE event, a speed improvement of nearly 50,000. Event generation became as fast as it was when the 1D generator was being used. Furthermore, because of the approach taken, this major speed improvement comes at no sacrifice in accuracy. Figure 4.5 shows a comparison of the average number of photoelectrons (npe) per event per PMT between water-phase SMELLIE data and simulations with both the old and new 2D generator. One can see clearly that both generators are as accurate as one another. Note that this plot uses the updated beam profiles as explained in the next section. 9 10 11 12 13 14 15 16 17 18

4.2 Improving the beam profiles

Even with the new 2D profile generator, a problem remains: the simulation fails to reasonably recreate data, and much of this appears to be because of the poor beam profile data being used. The curious “spot” for one of the fibres was already noted in the previous section that doesn't seem to be physical, and more broadly at large angles 20 21 22 23

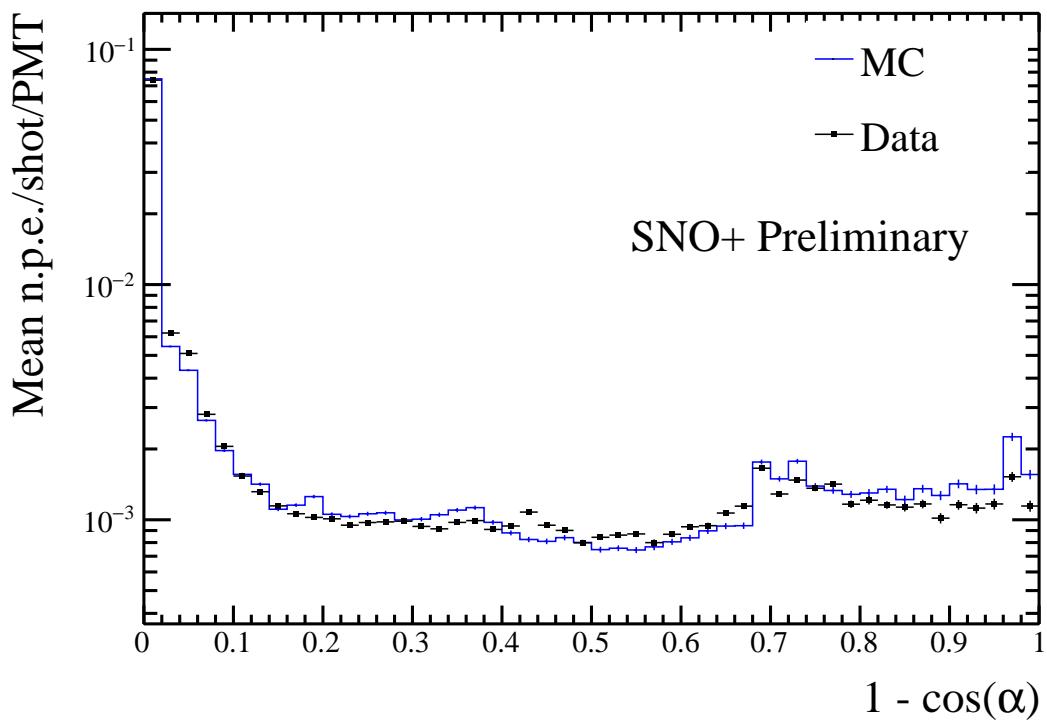


Fig. 4.5: Comparison of water-phase data to MC generated using both the old and new 2D beam profile generator approaches, with the updated beam profiles. Both versions of the generator are consistent with one another, but the new generator is many times faster.

 4.2 Improving the beam profiles

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for all the fibres there are large swathes of PMTs with an intensity of zero, providing
1 little useful information about the beam shape. It was shown in [5] that with the old
2 2D generator, the systematic uncertainty on the beam profiles was the dominant source
3 of error in the main SMELLIE analysis. To help improve this situation, it was decided
4 to update the existing beam profiles.
5

These old beam profiles were originally determined by looking at SMELLIE data
6 taken during the water-phase. Specifically, a “medium”-intensity subrun with one
7 of the lasers firing at a wavelength of 495 nm, was chosen for each fibre. “Medium”-
8 intensity corresponds to firing the relevant laser at a set intensity determined during
9 an earlier commissioning process, for which the maximum occupancy of PMT hits at
10 that intensity, i.e. the proportion of hits per event, corresponded to roughly 80%. This
11 value was chosen as it allowed for high statistics in a relatively short run-time, but not
12 so intense that the occupancy of any given PMT in the beamspot was 100%. Because
13 Rayleigh scattering is strongly-dependent on wavelength, the long wavelength of light
14 was chosen so that impacts from this scattering were small in the data.
15

SNO+ PMTs are unable to distinguish the exact number of photoelectrons being
16 generated. One is typically only able to know if a PMT has been triggered at all, by
17 any number of photoelectrons. As a result, the occupancy of a PMT over a number
18 of SMELLIE events, o , is a biased estimator of the mean number of photoelectrons
19 generated, μ . Assuming the number of photoelectrons generated in a given event
20 follows Poisson statistics, the probability of generating k photoelectrons is:
21

$$P(k|\mu) = \frac{\mu^k e^{-\mu}}{k!}. \quad (4.1) \quad 22$$

The probability of observing a “hit” in a given PMT corresponds to generating at
23 least one photoelectron:
24

$$P(\text{hit}|\mu) = P(k \geq 1|\mu) = 1 - P(k = 0|\mu) = 1 - e^{-\mu}, \quad (4.2)$$

which implies after rearrangement that one can determine the mean number of photoelectrons per event from the occupancy by:

$$\mu = \ln(1 - o). \quad (4.3)$$

This is the reason why we want to avoid PMTs with occupancies of 100%: they preclude one's ability to convert into a value for μ by looking at occupancy alone. We call this conversion from occupancy into npe the “multi-hit correction”. The impact of this correction is typically small for most PMTs, but can become very significant in a fibre's beamspot.

Once the npe mapping from data was obtained, a correction was then made for the detector's optics: even ignoring a fibre's beam profile, we still expect certain PMTs to be illuminated more than others because of e.g. reflections off the AV, or the solid angle subtended by the PMT bucket opening. For each fibre, a simulation was made where the beam profile was set as isotropic, and the corresponding npe mapping obtained: this map held information about the detector optics only. The beam profile mapping was then derived by simply dividing each fibre's npe mapping from data to its associated isotropic MC npe map. It is these maps that were first used in section 4.1.2.

4.2.1 Combining beam profile datasets

Fortunately, much more SMELLIE data was taken during the water-phase than was used for the original beam-profiling analysis. This additional data can be combined with that which was already used to far better constrain the beam profiles. In particular, given the existing assumption that scattering effects are minimal above wavelengths

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Run Number	Run Type	Comments
114,018	All PQ lasers; SuperK laser in 400–500 nm range	Only PQ495 laser and SuperK at 495 nm is used
114,023	SuperK laser in 500–600 nm range	Part 1 of this wavelength range; crash occurred on last subrun, so that subrun is ignored
114,034	SuperK laser in 500–600 nm range	Part 2 of this wavelength range

Table 4.1: Water-phase runs used for new beam profiling.

of ~ 490 nm, all data taken with wavelengths above this can also be used. The specific runs (and associated comments about their specifics) are described in Table 4.1. Because high-intensity runs require a different analysis approach (PMTs with high occupancies must use charge, not occupancy, to estimate npe), for this analysis we only considered subruns that used low or medium intensity set-points.

For each subrun j of data per fibre, we look only at PMT hits for each PMT i that has been identified as “good” for that subrun¹, $i \in G_j$. G_j here represents the set of good PMTs in subrun j . In particular, a “good” PMT must have valid electronic and timing calibrations, be at high voltage and masked into the detector’s trigger system for that subrun. In addition, an angular cut of $\alpha < 60^\circ$ was made to remove PMTs that are well outside any reasonable beam direction. To isolates the hits arriving directly from the fibre without reflecting, scattering, or being noise, a time cut was also made. Because what matters is the time relative to emission from the fibre, and the expected time-of-flight from fibre to different PMTs varies, a quantity known as the time residual was used. Starting with the calibrated hit time of a given PMT relative to the event’s trigger time, t_{hit} , the expected time-of-flight t_{TOF} from the fibre to the PMT was subtracted, estimated with the collaboration’s “Light Path Calculator”. Then, the emission time was also subtracted, t_{emm} , estimated by looking at the second-earliest value of $t_{hit} - t_{TOF}$ within the fibre’s central beamspot, defined as the PMTs for which

¹Strictly speaking, a PMT’s “goodness” is only determined on a run-by-run, not a subrun-by-subrun level, but this has no impact on the analysis.

¹ $\alpha < 3^\circ$. It was found that a “loose” time residual cut of $t_{res} \in [-10, +12]\text{ns}$ was
² sufficient to remove the vast majority of non-direct light with little signal sacrifice. In
³ the situation where a subrun with intensity was very small, it would not regularly have
⁴ at least two hits in the beamspot, and so the time residuals calculated would not be
⁵ valid for many events. To avoid this situation, a cut was made on any subruns with
⁶ mean intensities below 9 within their beamspot. This value was chosen as it would
⁷ mean a 2σ fluctuation downwards of $2 \cdot \sqrt{9} = 2 \cdot 3 = 6 \text{ npe}$ would still have more than
⁸ the 2 hits necessary for timing reconstruction. One fibre, FS207, has no data subruns
⁹ that satisfy this condition, and as such will have to be dealt with separately. For the
¹⁰ time being, this fibre was ignored.

¹¹ Extracting the underlying beam profiles from these data required some careful
¹² thought, especially because different subruns could have different intensities. Consider-
¹³ ing a PMT i in subrun j , the mean number of photoelectrons generated per event in
¹⁴ that PMT for that subrun, μ_{ij} can be decomposed as follows:

$$\mu_{ij} = I_j k_i = I_j b_i f_i. \quad (4.4)$$

¹⁵ I_j is the intensity of the subrun, i.e. the mean number of photons generated from the
¹⁶ fibre in that subrun per event. k_i is the probability that a given photon generated at
¹⁷ the fibre source ends up generating a photoelectron in PMT i . This itself can be further
¹⁸ split into two components: b_i , the probability that a given photon at the fibre source
¹⁹ points in the direction of PMT i ; and f_i , the probability that a given correctly-pointed
²⁰ photon actually makes it to the PMT and successfully generates a photoelectron. It is
²¹ b_i that is the actual beam profile we would like to measure.
²²

²³ Letting p_{ij} be the probability of observing a hit for a given event on a given
²⁴ PMT, the probability of observing m_{ij} hits out of N_j events in the subrun will be

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binomially-distributed:

$$P(m_{ij}|\mu_{ij}) = L(\mu_{ij}|m_{ij}) = \binom{N_j}{m_{ij}} p_{ij}^{m_{ij}} (1-p_{ij})^{N_j-m_{ij}} = \binom{N_j}{m_{ij}} (1-e^{-\mu_{ij}})^{m_{ij}} e^{-\mu_{ij}(N_j-m_{ij})}. \quad (4.5)$$

Here we have used equation 4.2, and noted that this probability distribution in m can be re-framed as a likelihood function for the parameter μ_{ij} . Considering only a single subrun of data, the maximum likelihood estimate of the parameter μ_{ij} can be shown to be:

$$\langle \mu_{ij} \rangle = -\ln \left(1 - \frac{m_{ij}}{N_j} \right) = \ln (1 - o_{ij}) \quad (m_{ij} \neq N_j), \quad (4.6)$$

where o_{ij} is just the occupancy of PMT i in subrun j . This is just the multi-hit correction formula seen in equation 4.3, which makes sense.

When looking at multiple subruns for the same fibre, the total likelihood function for a given PMT when considering all the data for a given fibre will be the product of the likelihoods from each dataset,

$$L(\{I_j\}, k_i | \{m_{ij}\}) = \prod_j L(I_j, k_i | m_{ij}) = \prod_j \binom{N_j}{m_{ij}} (1 - e^{-I_j k_i})^{m_{ij}} e^{-I_j k_i (N_j - m_{ij})}. \quad (4.7)$$

This leads to a log-likelihood distribution of

$$\mathcal{L}(\{I_j\}, k_i | \{m_{ij}\}) = \sum_j \left[\ln \left(\binom{N_j}{m_{ij}} \right) + m_{ij} \ln \left(1 - e^{-I_j k_i} \right) - I_j k_i (N_j - m_{ij}) \right]. \quad (4.8)$$

Formally, one could combine the likelihoods of all the PMTs together, and by looking at the maximum likelihood estimates for each of the parameters measure the parameter values this way. However, the set of equations one obtains through this approach quickly become analytically intractable, because the PMTs are coupled by the intensity values I_j . Even a direct numerical approach would be liable to fail: for a given fibre

¹ there can be dozens of subruns, and many thousands of PMTs of relevance, so the
² dimensionality of the system of equations would be far too large.

³ Because of this, a different approach was taken. It is expected that in a subrun the
⁴ total npe, summed over all good PMTs, should be proportional to the intensity value
⁵ I_j . One must be careful about this construction — different subruns can have different
⁶ sets of good PMTs, so two subruns with identical I_j values could have a larger summed
⁷ npe merely because more PMTs were good in that subrun. To counter-act this effect,
⁸ only PMTs that were classified as good in *all* subruns being analysed for that fibre
⁹ would be used for the npe summation. In other words, we use data from PMT i for
¹⁰ summing only if:

$$\text{11} \quad i \in \mathcal{I} = \bigcap_j G_j. \quad (4.9)$$

¹² We can then define the summed npe for a given subrun as $S_j = \sum_{i \in \mathcal{I}} \text{npe}_{ij}$, and assert
¹³ that $I_j = cS_j$. By finding a value proportional to I_j , there is now enough information
¹⁴ to maximise the log-likelihood $\mathcal{L}(k_i | \{m_{ij}\}, \{I_j\})$ with respect to k_i for each PMT
¹⁵ independently, and hence obtain estimates for these k_i parameters.

¹⁶ Of course, what is actually wanted are the underlying b_i values, not k_i . This is
¹⁷ where isotropic simulations come in. For each run of data used, a matching isotropic
¹⁸ MC was produced. As an example, a simulation for run 114,023 contained 200,000
¹⁹ events for each fibre using an isotropic beam profile, over the full wavelength range
²⁰ considered in this run, 500–600 nm, using the same run conditions as in data (which
²¹ PMTs were at high voltage, etc.).

²² For each isotropic MC run, both I_j^{MC} and k_i^{MC} were calculated via the method
²³ described above. Because the simulations were isotropic, the underlying value for b_i
²⁴ was constant across all the PMTs, and so $ak_i^{MC} = f_i$. By doing some rearranging of

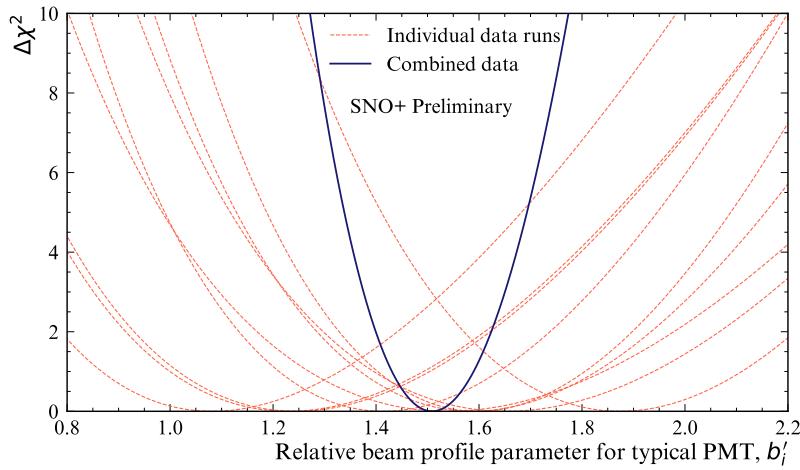


Fig. 4.6: Plot of $\Delta\chi^2 \simeq X_i$, twice the negative log-likelihood ratio, for both single subruns of a typical PMT, and when all relevant subruns are combined.

equation 4.4, we find that:

$$\mu_{ij} = I_j b_i f_i = c S_j b_i a k_i^{MC} = (acb_i) S_j k_i^{MC}. \quad (4.10)$$

As a result of this, given the set $\{S_j\}$ and k_i^{MC} , one can maximise the log-likelihood \mathcal{L} with respect to $b'_i = acb_i$ numerically, to obtain the maximum likelihood estimate of b'_i . Because a and c were global constants of proportionality, they would become irrelevant as soon as the beam profile was normalised in the CDF-creation process outlined in 4.1.2.

Figure 4.6 shows the shape of this log-likelihood distribution for a particular PMT when considering fibre FS007's beam profile. One can see how individual subruns provide much more information when combined than if one looked at a single subrun alone.

Another benefit of using this log-likelihood approach is that the resulting distribution's shape can be used for uncertainty estimation. In almost all cases, Wilks Theorem [6] allows us to produce 1σ confidence intervals about the maximum likelihood

estimate for b'_i , $\langle b'_i \rangle$, because

$$X(b'_i) = -2 [\mathcal{L}(b'_i) - \mathcal{L}(\langle b'_i \rangle)]$$

¹ approximates a χ^2 -distribution. As a result, the error bounds on our parameter estimate
² are given by when $X = 1$. The fact that the shape of X can be well-approximated by
³ a quadratic in the region near $X = 0$ indicates the validity of Wilks' Theorem being
⁴ used here.

⁵ Only a couple of exceptions to this approach of parameter estimation are possible.
⁶ In the case where $m_{ij} = N_j$, i.e. a PMT has 100% occupancy, no maximum likelihood
⁷ estimate exists: we need not worry about this, as subruns where this occurs have not
⁸ been used. On the other end, however, there are some PMTs for certain fibres where
⁹ after all subruns of data have been included, there remains no hits. In this scenario,
¹⁰ one can show that the log-likelihood becomes linear in the beam profile parameter:

$$\mathcal{L}(b'_i | \{m_{ij} = 0\}) = b'_i k_i^{MC} \cdot \sum_j [I_j N_j]. \quad (4.11)$$

¹² This scenario is very much reminiscent of rare-decay searches, and a similar approach
¹³ can be used. A 1σ upper limit on the possible value for b'_i can be analytically-calculated
¹⁴ to be:

$$b'_{i,ulim} = -\frac{k_i^{MC} \sum_j [I_j N_j]}{\ln [1 - \text{erf}(1/\sqrt{2})]}, \quad (4.12)$$

¹⁶ where $\text{erf}(x)$ is the error function. [18 pages for above two sections]

¹⁷ 4.2.2 Results & Discussion

¹⁸ WARNING: contents of this subsection will be gutted, focusing merely on impact of
¹⁹ combining data sets. Details about discrepancies will be covered in the next section.

4.2 Improving the beam profiles

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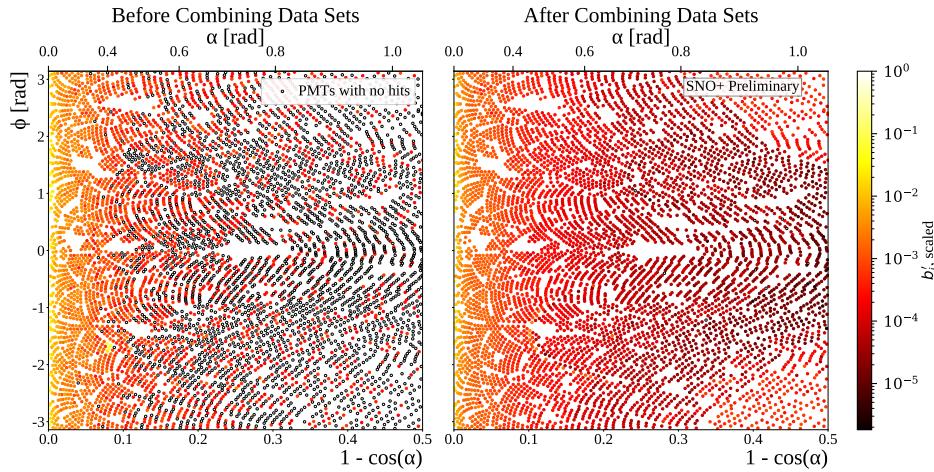


Fig. 4.7: Comparison between old and updated beam profiles for fibre FS055, after combining multiple data sets. Once again, the relative intensities (b'_i) for each PMT are given by the colour of each point, the position of each plotted in the 2D (r, ϕ) -space. The relative intensities have been both scaled here so that the largest value equals 1. Hollowed-out points are PMTs that, even after all relevant subruns have been combined, have no PMT hits.

[Will just be 1 page] Figure 4.7 shows the impact of using additional subruns of data on a typical beam profile. One can clearly see the great reduction in the number of PMTs with no hits in data. That many more data sets were included allowed for the major increase in dynamic range available for measuring these b'_i values. One can also note that by including additional data the curious spot that was seen in the old beam profile our at $r \approx 0.08$ has gone, further indicating that it was an artefact of that single data set.

Further details can be gathered from the interpolated intensity maps, one of which can be seen in figure 4.8. There are two curious stand-out features that can be seen here: firstly, there are multiple distinct parabolic arcs. These correspond to the shadows of the ropes that hold up/down the AV. More precisely, they are the mismodelling of those shadows — if the shadows were in the right place in the isotropic MC, then they would correctly cancel out any decreased intensity seen in the data of shadowed PMTs. These shadows could be mismodelled either because the positions of the ropes in the MC are in the wrong place, or the fibre's emission position is wrong. Note that any

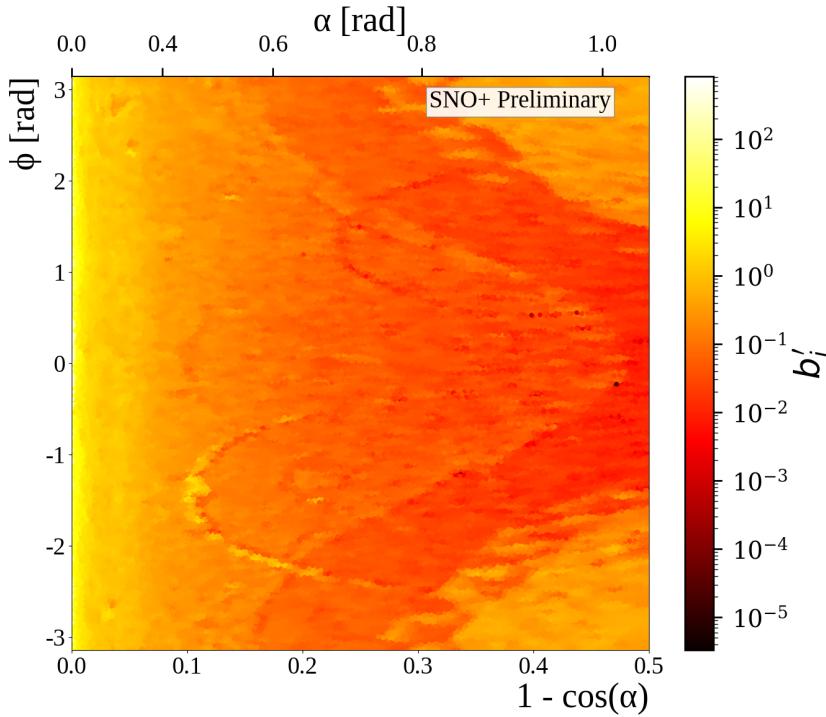


Fig. 4.8: Interpolated intensity map for the new updated beam profile of fibre FS055. The misalignment of rope shadows and AV effects, can both be seen.

1 mismodelling of the fibre's nominal emission direction has no impact on this shadowing
 2 problem, as changing that direction merely causes a change of basis in the (r, ϕ) -space.
 3 The latter possibility of incorrect fibre positions are more likely, and in fact these arcs
 4 in the beam profiles could be used as an effective way to correct for this problem.

5 The second distinctive feature of this intensity map is the large band of lower
 6 intensity varying between $r \approx 0.2 - 0.5$, followed by larger intensity out at large
 7 r values. This feature comes from light reflecting off the AV surface, or internally-
 8 reflecting. The reason for this band's functional dependence on ϕ is that this particular
 9 fibre, FS055, has a nominal fibre direction $\sim 10^\circ$ from pointing radially-towards the
 10 detector's centre. This feature appears in the updated beam profiles of all fibres, but
 11 its shape depends on the particular fibre's direction — for fibres pointing directly
 12 towards the detector's centre, there is little ϕ -dependence observed. Like the ropes,

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this feature must come from some form of mismodelling of the optics of the AV. A de-facto shadowing of PMTs in line with tangents from the AV surface which intersect the fibre position is to be expected. One also expects PMTs at polar angles larger than this to have their observed intensities boosted from reflected light off the AV. However, the discontinuities seen in the beam profiles indicate that for whatever reason this effect has been over-emphasised in the simulation.

There is a further phenomenon that can be seen, by comparing beam profile values obtained from a single subrun to the updated combined beam profile. This can be done by calculating the residuals corresponding to the single subrun, relative to the combined data set. The residual is negative if the combined data sets have a b'_i below the equivalent for a given single subrun; that is, the combined model underestimates this subrun for that PMT.

This information was plotted for two different subruns from the same fibre, seen in figure 4.9. One subrun was the same one used by Esther Turner for the original 2D beam profiling, with a wavelength of 495 nm; the latter was at the longer wavelength of 595 nm. For both subruns, most PMTs are seen to have intensities well-modelled by the combined model. However, there appears to be a significant amount of mismodelling within the beamspot. There also appears to be some systematic shift between data and model at somewhat larger polar angles. Moreover, this mismodelling seems not to be merely random, but a function of wavelength: at shorter wavelengths the beamspot tends towards being overestimated and then underestimated at larger values of α . At longer wavelengths, the beamspot becomes underestimated, with larger angles getting overestimated. This indicates that there appears to be a wavelength-dependence on the beam profiles, contradicting one of the main assumptions which we used to combine the water-phase data in the first place! All three of these features — rope shadows, AV reflections, and wavelength dependence — add systematic uncertainty to the beam

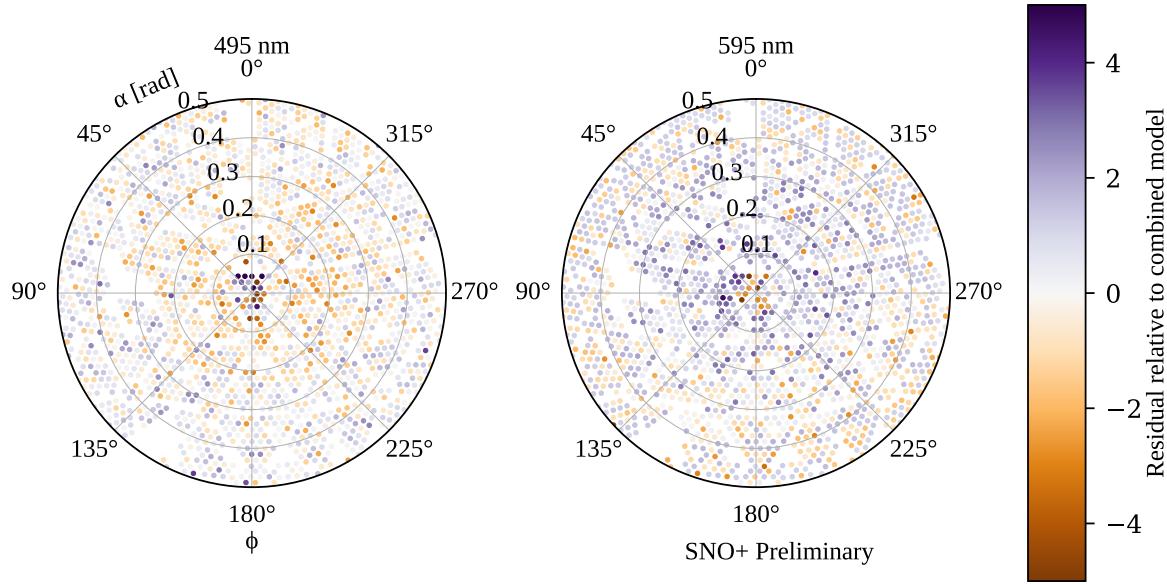


Fig. 4.9: Residuals from subruns at two different wavelengths, both compared to the combined beam profile model for fibre FS055. A negative sign, and hence bluer colours, indicate that the combined model underestimates the observed intensity for that particular subrun. Values with a magnitude beyond 5 are shown capped at this maximal value for the purposes of this plot. These PMTs are plotted in the polar fibre coordinates (α, ϕ) .

1 profiles, beyond the statistical uncertainty as measured by the width of the likelihood
 2 distribution. Certainly if one wanted to further improve the uncertainties in the beam
 3 profiles, tackling these challenges would be key.

4 4.3 Comparisons between Data and Simulation

- 5 • This focuses on disagreements noticed between data and MC even after the new
 6 generator and beam profiles have been used.
- 7 • Important to mention that none of these are necessarily game-ending, they are
 8 just systematics that may or may not be substantial in a given analysis with
 9 SMELLIE..

4.3.1 Forward Hemisphere Discrepancies

- The continued disagreement between data and MC when it comes to measured npe in various parts of the “forward” hemisphere. This includes:
 - The central beamspot,
 - The TIR region,
 - Rope shadows,
 - A noticeable wavelength-dependence.
- This is pretty much most of the contents of Section 4.2.2.

[4 pages]

9

4.3.2 Emission Time Discrepancies

- For certain lasers, a strong mismatch in the observed hit time residuals for prompt light.
- A mysterious +18 ns bump seen for the PQ495 laser.
- A trigger jitter in the SuperK laser.

[3 pages]

15

4.3.3 Backward Hemisphere Discrepancies

- The observed distribution of hits vs time and angle in MC does not match data in a number of ways for PMTs near the fibre emission point.
- Includes the outer-water scattering length, rope reflections, and investigations into whether certain modifications to the optics could plausibly fix things (so far, no).

21

¹ [6 pages]

² [32 PAGES TOTAL]

Chapter 5

Analysis of SMELLIE Data in the Scintillator Phase

This chapter contains two sets of analyses: measurements of the extinction lengths of the scintillator as a function of wavelength and time, as well as monitoring the Rayleigh scattering length over time.

5.1 Extinction Length Analysis

5.1.1 Motivation

- Explain motivating observations for this analysis: a substantial discrepancy between MC and data seen in the radial profiles of nhitsCleaned for ^{210}Po after the PPO top-up campaign. Performed by Serena.
- Hypothesis of a shortening of the absorption/scattering length proposed, further strengthened by Ben Tam’s ex-situ absorption measurements with scintillator taken from the detector, as well as knowledge about a likely “cooking” of PPO during the PPO-fill.

62

Analysis of SMELLIE Data in the Scintillator Phase

- 1 • Describe the provisional optics model decided on based on these measurements,
2 which includes an additional non-re-emitting component of the scintillator.
- 3 • As a further cross-check, SMELLIE should be sensitive to changes in the overall
4 extinction length of the scintillator, especially for short extinction lengths relative
5 to the size of the detector.
- 6 • More straightforward in measuring extinction length compared to scattering
7 length — no need to distinguish between scintillator re-emission and scattering.
- 8 • Further uses: can be used to monitor the extinction length over time!

9 [4 pages]

10 **5.1.2 Analysis Approach**

11 **The 2-PMT Case**

- 12 • Outline theoretical approach for how one could measure the extinction length of
13 scintillator through a comparison of SMELLIE data between the scintillator and
14 water phases, in the simplified 1 dimensional case with only 2 PMTs.

15 [3 pages]

16 **Combining Results Between PMTs**

- 17 • Not doing analysis with just 2 PMTs, of course! Can combine results from
18 multiple PMTs within a beamspot: I explain how here.

19 [2 pages]

5.1 Extinction Length Analysis**63****Corrections Between the Water and Scintillator Phases**

1

- Note the complications that we have to deal with. Namely, the differing refractive indices of the media bending the beamspot differently in the phases, as well as the method used to estimate t_{emm} .
2
3
4
- Explain how we deal with these, the former through MC simulation.
5

[2 pages]

6

5.1.3 Validation of the Analysis in Simulation

7

- Show results of this approach being used to measure the extinction length in simulation. How well does it do?
8
9

[3 pages]

10

5.1.4 Results in Data

11

- Describe the data used in this analysis, both water and scintillator, which can be shown in a table.
12
13
- Show examples of analysis of data in action for 375 nm data: typical t_{res} distributions of backscattered and beamspot PMTs; calculation of that particular extinction length measurement, followed by the graph for extinction length in 375 nm over all fibres and time periods.
14
15
16
17
- Discuss what results can be seen in this plot: consistency between fibres, the expected change as a function of PPO concentration, and stability of the extinction length during the main 2.2 g/L scintillator phase.
18
19
20

- Compare results to those made by Ben ex-situ: are they in agreement? If not, what possible systematics could there be? The main one for my analysis is likely to be uncertainties in the simulated beam profile that leak through into the refractive index correction of the beamspot. For the ex-situ analysis, the value of the extinction length obtained is achieved through background subtraction at some long wavelength, and the particular choice of this wavelength can lead to systematic changes in the obtained extinction length.
- Look at results at longer wavelengths: can anything reasonably be said at these longer wavelengths? Why/why not?
- Finally: describe any conclusions that can be reached, in particular whether we can affirm the optics model we use in RAT.

[8 pages]

5.2 Scattering Analysis

5.2.1 Historical Approaches and the Problem of Systematics

- Comparison to MC is necessary in scattering analysis, compared to merely being needed as a correction factor. This is because of the angular dependence of scattering. As a result, we can be far more susceptible to systematics from poor modelling!
- As a warning, show how Krish's/Esther's approach to the SMELLIE scattering analysis suffers majorly from these systematic effects. Requires describing their analysis approach briefly, and then explaining how the systematics described in Section 4.3 lead to major problems with this approach.

5.2 Scattering Analysis**65**

- Motivates the need for either reduced systematics, or an alternative analysis approach that is more robust to them! 1
2

[2 pages]

3

5.2.2 New Methodology

4

Signal Region Selection

5

- Propose the new analysis approach: looking at light in the “bad light-path” PMT region. Define what this region is. 6
7
- Give qualitative argument for why we expect this region to be robust to the beam profile systematics: dominated by the scattered signal as no direct light can make it here, and changes to beam profile should get “smeared out” after scattering. 8
9
10
11
- Show how simulations indicate this should be a region with a very high purity of scattered light, and (assuming all else being equal) robust to beam profile uncertainties. 12
13
14
- Confirm robustness of selected PMT region to uncertainties in AV offset and fibre position. 15
16

[5 pages]

17

Measuring the Emission Intensity

18

- Remaining systematics is now in the calculation of an average absolute emission intensity. 19
20

66

Analysis of SMELLIE Data in the Scintillator Phase

- 1 • Show how various methods don't work particularly well: whole detector npe,
- 2 beamspot npe, backscattered light npe, "bad light-path" PMTs but at later times.
- 3 Explain why it goes wrong for each method.
- 4 • Look at "beamspot but excepting the central bit": if that works well, then we
- 5 can continue!
- 6 • Otherwise, we'll have to live with measuring relative scattering lengths instead
- 7 of absolute amounts, using the outer water back-scattering as a measure of the
- 8 relative emission intensity.

9 [4 pages]

10 **5.2.3 Results**

- 11 • Actually do the proposed analysis on data, versus time and wavelength. Do the
- 12 results seem consistent between fibres? Are they sensible values?

13 [5 pages] [33 PAGES TOTAL]

Chapter 6

1

Solar Oscillation Analysis

2

Driving out into the Sun

Let the ultraviolet cover me up

Looking for a Creation Myth

Ended up with a pair of black lips

3

This is the End

PHOEBE BRIDGERS

Measuring the “solar” neutrino oscillation parameters Δm_{21}^2 and θ_{12} is one of the principal aims of the SNO+ detector during the scintillator-phase. There are, in fact, two complementary methods of measuring these parameters: the oscillations of anti-neutrinos from terrestrial nuclear reactors, and the oscillations of neutrinos from the Sun.

4

5

6

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8

This chapter focuses on the latter approach, using ${}^8\text{B}$ neutrinos coming from the Sun to measure the solar oscillation parameters. An initial background-free study was performed by Javi Caravaca [7], which demonstrated that it was indeed possible to make such a measurement in the detector. The work in this chapter builds on substantially from that analysis. This chapter also draws on the associated reactor anti-neutrino

9

10

11

12

13

¹ analysis built by Iwan Morton-Blake [3], and more broadly from the general techniques
² used in the $0\nu\beta\beta$ analysis of Tereza Kroupova [8] and Jack Dunger [9].

³ This chapter begins by explaining how it is possible to measure the solar oscillation
⁴ parameters via ^8B events. Then, the framework used to perform the analysis is then
⁵ explained: that of a *Bayesian Analysis using Markov Chain Monte Carlo techniques*.
⁶ After the method has been described, the dataset upon which the analysis is performed
⁷ is then introduced. The results and associated validation are then given. Given these
⁸ results, a projection is then made for the expected sensitivity to θ_{12} as a function of
⁹ livetime.

¹⁰ 6.1 Analysis Methodology

¹¹ 6.1.1 Observational Principle

¹² How can we measure neutrino oscillation parameters via solar neutrinos in the SNO+
¹³ detector? As discussed in Chapter 1, it is possible to detect all flavours of neutrino
¹⁴ through elastic scattering with electrons in the detector. If this interaction was purely
¹⁵ neutral-current, then there would be no way of telling the flavour-state of an interacting
¹⁶ neutrino. However, electron neutrinos are able to interact through an additional
¹⁷ charged-current mode. This modifies the cross-section for electron neutrinos, and
¹⁸ means that as the survival probability for electron neutrinos generated from the Sun,
¹⁹ P_{ee} , is modified, the interaction probability of neutrinos with the detector will also.

²⁰ Of course, we do not directly measure neutrino energies in the detector — only
²¹ the associated scattered electron. If there were no correlation between the observed
²² electron energy and its associated neutrino, then the only effect of neutrino oscillations
²³ would be to change the overall observed rate of events due to this process. There
²⁴ would be no change in the shape of the event’s energy spectrum, even though neutrino

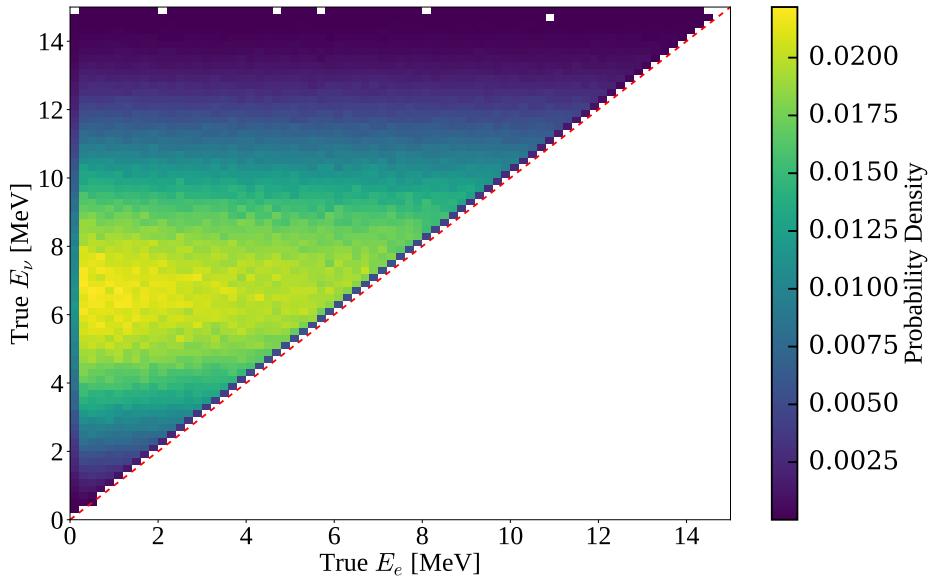


Fig. 6.1: 2D probability distribution comparing the true neutrino energy from a ${}^8\text{B}$ ν_e to the true energy of the scattered electron. Also shown in red is the line $E_\nu = E_e$.

oscillations are a function of neutrino energy. Fortunately, there is some dependence of
1
the neutrino's energy, E_ν , on that of the scattered electron, E_e . This dependence can
2
be seen in Fig. 6.1 for ${}^8\text{B}$ electron neutrinos interacting in SNO+. As can be seen, the
3
dependence is weak, and comes mostly from basic energy conservation: If one observes
4
a 10 MeV electron event in the detector, it can't reasonably have come from a 5 MeV
5
neutrino.
6

In Fig. 6.2 we can see the impact each physical process has on the energy spectrum
7
that we eventually observe. We start with a broad energy distribution of ${}^8\text{B}$ electron
8
neutrinos generated in the Sun. These neutrinos then oscillate their flavour state as
9
they propagate to the detector, in an energy-dependent manner. When a (tiny) fraction
10
of these neutrinos interact with the electrons in our detector, there is both an energy-
11
and flavour-dependence on the cross-section. The scattered electrons gain a kinetic
12
energy with some mild dependence on the inciting neutrino's energy, which is then
13
measured by the detector to within some energy resolution.
14

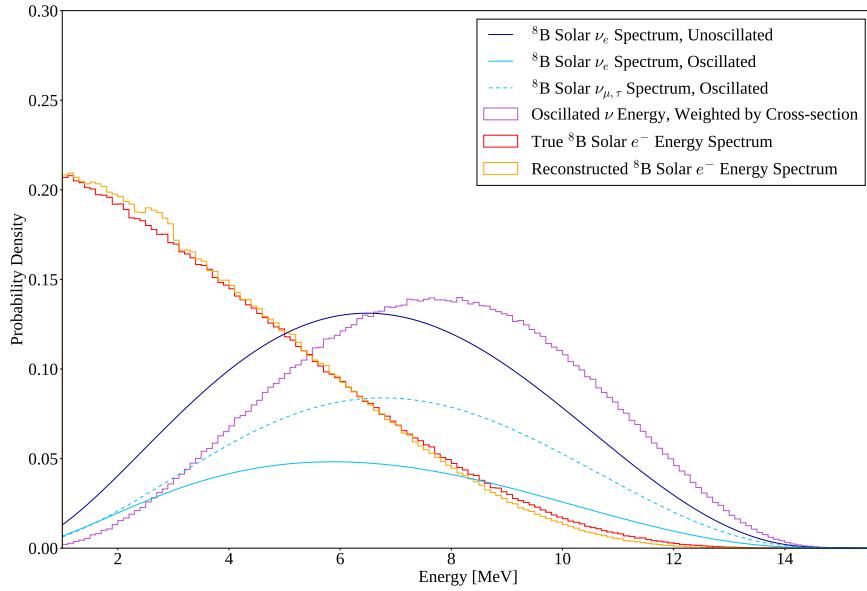


Fig. 6.2: The evolution of energy distributions related to ${}^8\text{B}$ solar neutrino detection. The unoscillated neutrino spectrum is taken from [10]; neutrino oscillations assume oscillation parameters from the current global fit results [11] and ν_e survival probabilities calculated via the method described in Section 6.1.8. The latter three distributions were obtained from MC production as described in Section 6.2.1, with the cross-section formula coming from [12].

Fig. 6.3

Let us now consider the dependence of P_{ee} on the individual neutrino oscillation parameters. Recall from Eq. 1.1 that, after considering matter-induced oscillations due to neutrinos passing through the Sun and possibly the Earth, $P_{ee} = P_{ee}(\tan 2\theta_{12}^M, \sin \theta_{13}^M, \Delta m_{21,M}^2) = P_{ee}(\theta_{12}, \theta_{13}, \Delta m_{12}^2, \Delta m_{13}^2)$. Fig. 6.3 shows the dependence of each of these four oscillation parameters on $P_{ee}(E)$. We can see that in reality only the two parameters Δm_{21}^2 and θ_{12} have a substantial impact on $P_{ee}(E)$ and hence the observed electron energy spectrum. Because of this, for this analysis we will only ever vary these two oscillation parameters, and keep θ_{13} and Δm_{13}^2 at their current global fit values¹ of $\sin^2 \theta_{13} = 0.0222$ and $\Delta m_{13}^2 = +2.515 \times 10^{-3} \text{ eV}^2$ [11].

¹We use the global fit results excluding Super-Kamiokande’s atmospheric data, and assuming normal ordering of the neutrino mass hierarchy. This choice has a tiny impact on the magnitudes of these two fixed parameters, the main impact being the sign of Δm_{13}^2 .

6.1.2 Background Processes

Sadly, elastically-scattered electrons from ^8B neutrinos are not the only events we see in the SNO+ detector during the scintillator phase. There are a number of background processes that our signal must compete against. Below a reconstructed energy of $\sim 2.5 \text{ MeV}$, it is known that various backgrounds completely swamp any possible ^8B signal, and so for this analysis we only consider processes that can generate reconstructed energies of at least $E_{\min} = 2.5 \text{ MeV}$. The following subsections explain each of these backgrounds, as well as methods that have been used to mitigate them as much as possible.

Internal Uranium- and Thorium-Chain Backgrounds

Although every effort has been made to make the scintillator cocktail that fills SNO+ to be as radio-pure as possible, there inevitably remain trace amounts of the radioactive isotopes that derive from the decay chains of the ^{238}U and ^{232}Th isotopes. Fig. 6.4 shows these two decay chains. Fortunately, only a fraction of the radioactive isotopes in these chains actually are capable of generating events in the detector with energies above E_{\min} : these have been highlighted in Fig. 6.4 in gold.

Of particular note are the decays of ^{212}Bi and ^{214}Bi . Both are capable of either α - or β -decays to Tl or Po isotopes, respectively. For the former, it is the subsequent β -decay of the Tl that can have a reconstructed energy above E_{\min} . For the latter, the Bi decay is the part of the pair of decays that can lie above E_{\min} . Although the α -decays here certainly have Q-values well above 2.5 MeV, the liquid scintillator quenches the observed energy well below E_{\min} . The so-called “Bi–Po” decays are particularly special because the lifetimes of ^{212}Po and ^{214}Po are 300 ns and 164 μs , respectively, which are short enough to allow for highly-effective coincidence tagging.

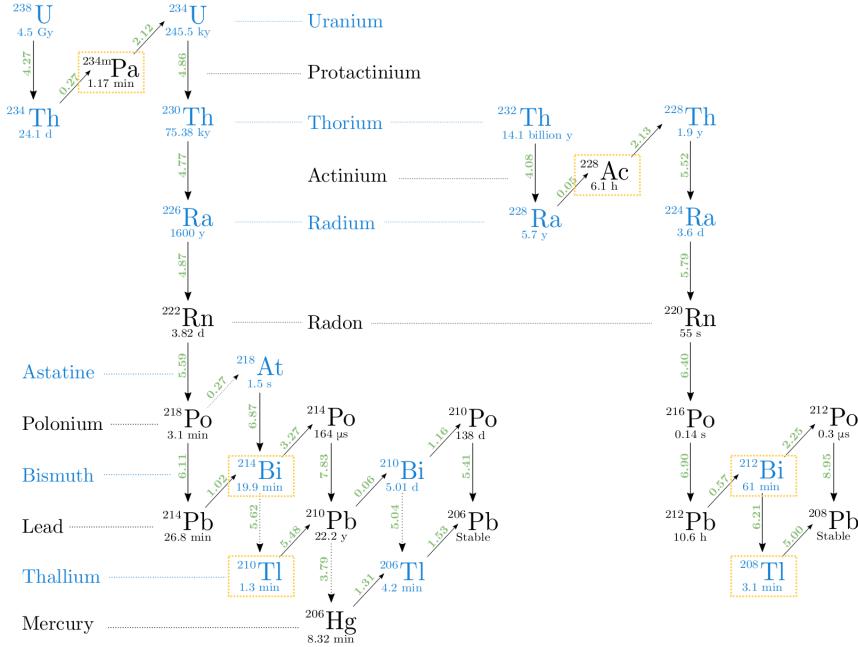


Fig. 6.4: The ^{238}U and ^{232}Th decay chains, taken from [8]. Isotopic half-lives are given below their symbol; the Q-values for each decay, in MeV, is given in green. Downward arrows indicate an α -decay; diagonal arrows indicate β -decay. Isotopes highlighted in gold are potential backgrounds for this solar analysis.

There are two classes of Bi–Po event in the detector: “out-of-window” events for which the Bi and Po occur in separate event windows, and “in-window” events whereby the Bi and Po occur within the same event window. These lead to two distinct strategies for tagging these kinds of events. For out-of-window Bi–Pos, we look for a delayed coincidence of two events. Using the tagging algorithm suggested in [] as a starting point, the chosen procedure was as follows. There must be two events that trigger the detector within $4\ \mu\text{s}$ of one another, and both have a valid `scintFit` position reconstruction within $2\ \text{m}$ of one another. The delayed candidate event must also have at least 100 cleaned PMT hits. This very broad coincidence tagging procedure was designed to ensure that the cut was as *efficient* in tagging (and hence, rejecting) Bi–Pos as possible, whilst negligibly impacting the solar signal. This is in contrast to the cuts chosen by Rafael Hunt-Stokes in [], which try and obtain a highly *pure* sample of Bi–Po tags.

Fig. 6.5

Of course, the above delayed coincidence procedure cannot catch any of the in-window Bi–Po events. For these, we use a different approach. Because two decays happened in the same event window, we expect to see two distinct peaks in the event’s time residual spectrum. In order to look for this event topology, a likelihood-ratio classifier was run over events, first developed by Eric Marzec [] and re-coordinated for the 2.2 g l^{-1} LABPPO scintillator optics by Ziping Ye []. This classifier calculates the likelihood ratio between the null hypothesis of a $0\nu\beta\beta$ event (a proxy in this analysis for single-site events such as our ${}^8\text{B}$ signal) and the alternative hypothesis of an in-window Bi–Po event. The more negative the value of the result, `alphabeta212`, the greater the evidence there is for rejecting the null hypothesis of a single-site event. Events with $\text{alphabeta212} < 0$ were then rejected.

Combining both out-of-window and in-window Bi–Po tagging, the impact on ${}^{212}\text{Bi–Po}$, ${}^{214}\text{Bi–Po}$, and ${}^8\text{B } \nu_e$ events can be seen in Fig. 6.5. We consider here only events that pass all other cuts used in this analysis: see Section 6.2.2 for the specifics of the cuts used. Because of the different lifetimes of the decays, ${}^{214}\text{Bi–Po}$ decays predominantly fall out-of-window whilst ${}^{212}\text{Bi–Po}$ events are typically in-window. This explains why the out-of-window tagging is substantially better at cutting ${}^{214}\text{Bi–Po}$ decays, whereas the in-window tagging far better tags ${}^{212}\text{Bi–Po}$ decays. Overall, within the analysis region of interest (ROI), the two combined cuts are able to tag TODO% of ${}^{214}\text{Bi–Po}$ triggered events, TODO% of ${}^{212}\text{Bi–Po}$ triggered events, whilst retaining TODO% of ${}^8\text{B } \nu_e$ signal events.

(α, n) Reactions

The impact of ${}^{238}\text{U}$ - and ${}^{232}\text{Th}$ -chain isotopes does not simply end at their direct decays. It is possible for the α s generated during these decays to undergo their own interactions

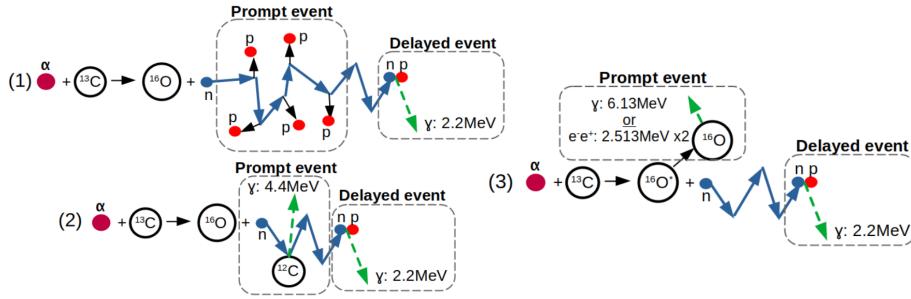


Fig. 6.6: Schematic of the three dominant modes of (α, n) interaction, taken from [3].

Fig. 6.7

with nuclei in the detector. Within the organic scintillator of SNO+, the dominant interaction of this type is when an α collides with a ${}^{13}\text{C}$ nucleus, emitting a neutron: $\alpha + {}^{13}\text{C} \longrightarrow {}^{16}\text{O} + n$. This is known as an (α, n) reaction.

The topology of this reaction in the detector is a delayed coincidence, as shown in Fig. 6.6. For the prompt signal, there is the light emitted from the α just before, and the n just after the (α, n) . The neutron then thermalises and gets captured by another nucleus — usually hydrogen in SNO+ — which creates an excited state that then eventually decays, creating a γ that creates the delayed signal in the detector [].

As can be seen in Fig. 6.7, (α, n) interactions can lead to events reconstructed at a wide variety of energies, which could be an issue for this analysis. However, because they are delayed coincidence events with a typical decay time of $\sim 100\text{ ns}$, the aforementioned out-of-window Bi–Po tagging algorithm also efficiently tags (α, n) events. Looking again at Fig. 6.7, simply by using the out-of-window Bi–Po tagger without any further modifications TODO% of events in the ROI are cut.

15 External Backgrounds

All materials within the SNO+ detector are radioactive, not just the liquid scintillator cocktail. This includes the acrylic, ropes, external water, and PMTs. These components

Fig. 6.8

have had their radiopurity “assayed” (that is, measured) throughout the detector’s lifetime, often back to the construction of the original SNO detector itself. The materials other than the liquid scintillator are known to have far higher background levels, especially in the important ^{238}U - and ^{232}Th -chain backgrounds [13]. To distinguish between the inherent backgrounds within the scintillator, and the backgrounds from materials at larger radii, we use the terminology “internal” and “external”, respectively.

Although there are numerous external backgrounds, with a suitably accurate and precise position reconstruction algorithm they can be suitably handled. The simplest approach is with a so-called “fiducial volume” (FV) cut: just throw out all events that reconstruct beyond some radius. The only external background events that will reach within the FV are those that have reconstructed very poorly, or have some long-distance radiation that manages to deposit radiation close to the centre of the AV. Because α and β radiation can only travel short distances through the detector, it is only γ radiation that can realistically travel far enough into the detector to be able to reconstruct anywhere near the centre. Moreover, the intensity of this γ radiation attenuates exponentially towards the centre of the detector, meaning only a tiny fraction of the total number of external events reconstruct within a 3.5 m FV, say. This strong radial-dependence can be seen in Fig. 6.8.

What this figure also demonstrates is that our solar signal has a completely different radial dependence to these backgrounds. As a result, if one considers not just the energy of events but also their reconstructed radius, then it is possible to get an additional handle on the external backgrounds. The FV cut can then be pushed further out to larger radii, allowing one to gain more signal statistics.

Work by Tereza Kroupova [8] allows for additional means of distinguishing external backgrounds from the solar signal. The underlying assumption in the reconstruction of

Fig. 6.9

¹ SNO+ events is that there was an electron at a single point, which is entirely valid for
² ^8B elastic scattering events. However, external backgrounds can fail this assumption in
³ two ways. Firstly, these radioactive decays often generate γ radiation in addition to the
⁴ main α/β particle, which creates a multi-site event. Because the `scintFit` position
⁵ reconstruction algorithm is not prepared for a distribution of energy depositions in the
⁶ scintillator, the t_{res} distribution will broaden. This allows an event classifier to be built
⁷ that distinguishes between the t_{res} distributions of single-site events and externals,
⁸ known as the “external background timing classifier”. Secondly, because external
⁹ backgrounds that do reconstruct close to the centre of the detector typically have a γ
¹⁰ that travelled a long distance towards the centre of the detector, we expect the earliest
¹¹ light that hits the PMTs to arrive most often along the direction of the reconstructed
¹² position vector. A distribution of PMT hits for a given event as a function of their
¹³ angular distribution relative to the direction of position reconstruction can be built,
¹⁴ and compared to the expected distributions for single-site and external background
¹⁵ events. This is known as the “external background topological classifier”. Much like
¹⁶ the classifier described in Section 6.1.2, the single-site events used for comparison were
¹⁷ $0\nu\beta\beta$ events, but these have a similar single-site structure to the solar signal of interest
¹⁸ in this analysis.

¹⁹ Fig. 6.9 shows the correlation between the two classifier results for both a typical
²⁰ external background, and $^8\text{B} \nu_e$ events.

²¹ Cosmogenic Isotopes

²² The final source of background events are radioactive isotopes that form via collisions of
²³ cosmic rays with atomic nuclei, known as cosmogenic isotopes. Most of these isotopes
²⁴ are short-lived [], with lifetimes $\mathcal{O}(1\text{s})$. Fortunately, the depth of SNO+ means that

our rate of cosmic ray muons interacting with the detector is only 3 an hour [1]. Because the rate is so low relative to other experiments, relatively straightforward approaches to tagging and removing cosmic ray muons and their cosmogenic followers can be utilised without substantial loss of livetime. Events are tagged as a cosmic ray muon if they create a sufficient number of hits in the outward-looking PMTs above background levels, as well as many hits within the detector itself. The details of this tagging for the scintillator-fill were put in place by Lorna Nolan [2], modifying the existing algorithms used in the water-phase [3] and in SNO [4]. After a tagged cosmic muon event, all events for the next 20 s are then vetoed as a means of rejecting followers. This simple cut is enough to remove the vast majority of expected cosmogenic events in the scintillator-phase. The expected impact on loss of livetime, and hence quantity of signal events, is 3 lots of 20 s vetoes an hour, that is to say 1/60th of the signal is lost via this cut.

There is one cosmogenic isotope with a long-enough half-life that even the 20 s muon follower veto is not sufficient to remove all events. This is ¹¹C, which β^+ -decays to stable ¹¹B with a half-life of 20.4 min. The maximum possible energy deposited in this decay is 1.982 MeV [5], just below E_{\min} , so only a small fraction of ¹¹C events end up in the ROI: the ones with very high energies that get their energy reconstruction falsely-inflated by some amount. As a result, this background is expected to be very much sub-dominant to all other backgrounds in this analysis. Because this background is important to consider for some other analyses, a triple-coincidence tagging algorithm is being built by Katharine Dixon [6], but not used for this analysis currently.

6.1.3 The Log-likelihood Test Statistic

At the highest level, this analysis involves taking the data observed in the scintillator-fill after applying a certain set of cuts, along with simulated PDFs for all processes believed

1 to build up the observed data with those same cuts applied, and then attempting to fit
 2 the combined energy and radial distributions of the MC to that of the data. Given a
 3 set of PDFs, to try and match the distribution of observables in data we can modify
 4 a number of parameters. These consist of the normalisations of each PDF (i.e. the
 5 total number of events observed due to that process), and any systematic parameters
 6 that could modify the shapes of these distributions. For this analysis, the neutrino
 7 oscillation parameters act as *de facto* systematic parameters, as they modify the shape
 8 of the ${}^8\text{B}$ PDFs. Of course, unlike usual systematics the oscillation parameters are
 9 what we are actively trying to measure, instead of being a nuisance.

10 In order to perform a fit to data in this way, we must first answer a set of questions:

- 11 1. Which signal and background processes must we consider?
- 12 2. In addition to their normalisations, are there any further parameters necessary
13 to specify the distributions of the PDFs for each of the processes? Systematics
14 and oscillation parameters are good examples.
- 15 3. What is our test statistic?
- 16 4. What algorithm do we use to try and find the best-fit result?
- 17 5. How do we measure uncertainties on these best-fit values for each parameter?

18 In section 6.1.2, question 1 was answered for this analysis. We now give the answer to
 19 question 3; all other questions on this list will be answered shortly.

20 The test statistic used for this analysis is the *binned extended log-likelihood*. Once
 21 the data and MC PDFs have been binned in both the observables of interest, it is
 22 assumed that the expected number of events in a given bin j is governed by a Poisson
 23 distribution:

$$24 \quad P_j(n_j|\lambda_j) = \frac{\lambda_j^{n_j} e^{-\lambda_j}}{n_j!}, \quad (6.1)$$

where $P_j(n_j|\lambda_j)$ is the probability of observing n_j events in bin j , given an expectation of λ_j events in total from signal and background processes in that bin. This λ_j can be decomposed into each of the expected rates for each process, i :

$$\lambda_j = \sum_{i=1}^{N_{\text{PDFs}}} \mathcal{N}_i P_{ij}(\boldsymbol{\theta}), \quad (6.2)$$

where N_{PDF} is the number of PDFs being considered in the analysis, \mathcal{N}_i is the normalisation parameter of the i^{th} PDF, and $P_{ij}(\boldsymbol{\theta})$ is the probability of observing an event of process type j in bin i , assuming a set of non-normalisation parameters $\boldsymbol{\theta}$. By combining the probabilities of all the bins together, the total probability for a given set of processes assuming these parameters to give rise to the data seen is:

$$P(\mathbf{n}|\mathcal{N}, \boldsymbol{\theta}) = L(\mathcal{N}, \boldsymbol{\theta}|\mathbf{n}) = \prod_{j=1}^{N_{\text{bins}}} \frac{\left[\sum_{i=1}^{N_{\text{PDFs}}} \mathcal{N}_i P_{ij}(\boldsymbol{\theta}) \right]^{n_j} e^{-\sum_{i=1}^{N_{\text{PDFs}}} \mathcal{N}_i P_{ij}(\boldsymbol{\theta})}}{n_j!}, \quad (6.3)$$

where N_{bins} is the total number of bins being considered in the analysis. This probability can be re-framed as the likelihood of the vectors of parameters \mathcal{N} and $\boldsymbol{\theta}$ given the vector of number of events in each bin, \mathbf{n} : $L(\mathcal{N}, \boldsymbol{\theta}|\mathbf{n})$. It is rare to see the likelihood as-is, instead, for computational purposes the log-likelihood is used instead, $\mathcal{L}(\mathcal{N}, \boldsymbol{\theta}|\mathbf{n}) := \ln L(\mathcal{N}, \boldsymbol{\theta}|\mathbf{n})$. We can then get to the formula actually used for this analysis:

$$\mathcal{L}(\mathcal{N}, \boldsymbol{\theta}|\mathbf{n}) = - \sum_{i=1}^{N_{\text{PDFs}}} \mathcal{N}_i + \sum_{j=1}^{N_{\text{bins}}} n_j \ln \left(\sum_{i=1}^{N_{\text{PDFs}}} \mathcal{N}_i P_{ij}(\boldsymbol{\theta}) \right). \quad (6.4)$$

6.1.4 The Bayesian Statistical Approach

There are two main schools of statistical inference, “Frequentist” and “Bayesian”. In the former, probabilities describe the fraction of times a situation can be found within the whole ensemble of possible worlds. For the latter, we care not about an ensemble

¹ of worlds but instead our degree of belief in this current one. We update our beliefs as
² we acquire knowledge of the world through Bayes' Theorem:

$$\text{³} P(\boldsymbol{\mu}|\mathbf{x}) = \frac{\mathcal{L}(\boldsymbol{\mu}|\mathbf{x}) P(\boldsymbol{\mu})}{P(\mathbf{x})}. \quad (6.5)$$

⁴ Here, $\boldsymbol{\mu}$ is the set of parameters that model our system, $P(\boldsymbol{\mu})$ is our *prior* (pre-existing)
⁵ distribution for those model parameters, and \mathbf{x} is the data taken in our experiment. The
⁶ updated, *posterior* distribution $P(\boldsymbol{\mu}|\mathbf{x})$ is then the prior multiplied by the likelihood
⁷ of parameters $\boldsymbol{\mu}$ given observations \mathbf{x} , $L(\boldsymbol{\mu}|\mathbf{x})$, and divided by the total probability
⁸ $P(\mathbf{x})$ of observing \mathbf{x} under any circumstance.

⁹ Both approaches to statistics are widely-used in statistical analysis, in both particle
¹⁰ physics and beyond. The Bayesian approach was used for this analysis, as it was
¹¹ believed that this helps keep transparent what assumptions are being made in the
¹² analysis.

¹³ Now, if one is able to determine the overall posterior distribution, then it is possible
¹⁴ to derive best-fit values with uncertainties for all parameters in the fit. This is done by
¹⁵ “marginalising” the posterior distribution, i.e. integrating over all parameters other
¹⁶ than the one of interest. A sensible best-fit value is then the modal marginalised
¹⁷ posterior density, the highest value in the marginalised distribution. The uncertainty
¹⁸ on this value is derived from the spread of the marginalised posterior, by the calculation
¹⁹ of the 1σ Credible Interval (CI): this is the set of values for a given parameter which
²⁰ has a total posterior probability of 68.3%, and contain the best-fit value. There are in
²¹ fact an infinite number of CIs that satisfy this property; for this analysis, the values
²² are chosen in decreasing order of marginalised posterior probability density.

6.1.5 Markov Chain Monte Carlo

Of course, all of this assumes that one can accurately determine the posterior density distribution. Whilst the likelihood and prior distribution are straightforward enough to calculate, often-times $P(\mathbf{x})$ (which acts as a normalisation) is very challenging to determine. This is because calculating this normalisation involves integrating the likelihood over all the parameter space, and if there are a large number of parameters this can become enormously numerically complex. An alternative approach is needed!

That alternative comes in the form of *Markov Chain Monte Carlo*, MCMC. A Markov Chain is any mathematical system for which the next state of the system is dependent only on its current state; the system is in some sense “memoryless”. For a large class of Markov Chains — those that are “ergodic” and “aperiodic” — one can prove that regardless of the initial position on the chain, the probability distribution converges to the same distribution []. MCMC uses such a Markov Chain which attempts to converge towards the posterior density distribution in particular. In MCMC, after choosing the initial position in the parameter space, successive states are chosen at random with a probability dependent only on the properties of the current position in parameter space and the proposed position. The convergence property of Markov Chains means that the set of steps made in the parameter space after some initial number of steps will have a distribution that converges to that of the posterior density distribution.

There are a number of MCMC algorithms, and the particular one used in this analysis is that of the *Random-Walk Metropolis-Hastings Algorithm*. In this algorithm, after the initial position in the parameter space $\boldsymbol{\mu}$, a new step is proposed some distance from the current one, $\boldsymbol{\mu}'$. This step is chosen at random from a multivariate Gaussian distribution centred on the current position, with widths in each dimension of the parameter space chosen beforehand as constants for tuning the MCMC process. This

¹ choosing of a new proposed step at random is what gives the algorithm its Monte Carlo
² and Random Walk titles. Once a new step is proposed, it is accepted as the new step
³ with a probability $S(\boldsymbol{\mu}'|\boldsymbol{\mu})$ according to the condition of *detailed balance*:

$$\begin{aligned} \text{⁴} S(\boldsymbol{\mu}'|\boldsymbol{\mu}) &= \min\left(1, \frac{P(\boldsymbol{\mu}'|\mathbf{x})}{P(\boldsymbol{\mu}|\mathbf{x})} \frac{R(\boldsymbol{\mu}|\boldsymbol{\mu}')}{R(\boldsymbol{\mu}'|\boldsymbol{\mu})}\right) = \min\left(1, \frac{L(\boldsymbol{\mu}'|\mathbf{x}) P(\boldsymbol{\mu}')}{L(\boldsymbol{\mu}|\mathbf{x}) P(\boldsymbol{\mu})} \frac{R(\boldsymbol{\mu}|\boldsymbol{\mu}')}{R(\boldsymbol{\mu}'|\boldsymbol{\mu})}\right) \\ \text{⁵} &= \min\left(1, \frac{R(\boldsymbol{\mu}|\boldsymbol{\mu}')}{R(\boldsymbol{\mu}'|\boldsymbol{\mu})} \exp\left[\mathcal{L}(\boldsymbol{\mu}'|\mathbf{x}) - \mathcal{L}(\boldsymbol{\mu}|\mathbf{x}) + \ln \frac{P(\boldsymbol{\mu}')}{P(\boldsymbol{\mu})}\right]\right). \end{aligned} \quad (6.6)$$

⁷ $R(\boldsymbol{\mu}'|\boldsymbol{\mu})$ is the probability density that position $\boldsymbol{\mu}'$ is proposed as a step from position $\boldsymbol{\mu}$,
⁸ and vice versa for $R(\boldsymbol{\mu}|\boldsymbol{\mu}')$. In most cases, because of the use of the same multivariate
⁹ Gaussian in choosing proposals, $\frac{R(\boldsymbol{\mu}|\boldsymbol{\mu}')}{R(\boldsymbol{\mu}'|\boldsymbol{\mu})} = 1$ simply. This component only becomes
¹⁰ important at the edges of the parameter space, preventing the sampling probability
¹¹ being incorrectly impacted if a proposed step goes outside the allowed parameter space.

¹² It is the detailed balance condition that ensures convergence of the MCMC algorithm
¹³ to specifically the posterior density distribution. Crucially, because it is only dependent
¹⁴ on the ratio of posterior densities, the hard-to-calculate normalisation $P(\mathbf{x})$ in both
¹⁵ posterior density terms cancels out, meaning one only needs to calculate the likelihood
¹⁶ and priors for each step, as well as $\frac{R(\boldsymbol{\mu}|\boldsymbol{\mu}')}{R(\boldsymbol{\mu}'|\boldsymbol{\mu})}$.

¹⁷ The specific implementation of MCMC used for this analysis is that of **OXO**, a C++
¹⁸ analysis framework first developed by Jack Dunger [9]. **OXO** is able to run the Metropolis-
¹⁹ Hastings algorithm on multidimensional binned data, using the log-likelihood defined
²⁰ in 6.1.3. This framework also allows one to include systematic parameters that can
²¹ float within the fit, and define non-uniform priors for normalisations and systematics
²² that have constraints: the details of this will be discussed shortly.

6.1.6 Choosing Priors

For this analysis, the suggestions made by Biller & Oser in [14] about choosing prior distributions are followed: for parameters that do not have some pre-existing constraint, a flat prior is used. A nice consequence of this choice is that $\ln \frac{P(\mu')}{P(\mu)} = 0$, so the actual value of the prior for these variables never needs to be calculated when running the MCMC algorithm. For the bulk of this analysis, uniform priors are assumed on the neutrino oscillation parameters Δm_{21}^2 and θ_{12} , as the magnitudes of these parameters are now well-established.

For parameters with existing asymmetric constraints $\beta_{-\sigma_-}^{+\sigma_+}$, this analysis uses an asymmetric Gaussian prior, equivalent to the logarithm of the prior being an asymmetric quadratic:

$$\ln P(\mu) = \mathcal{A} - \begin{cases} \frac{(\mu-\beta)^2}{2\sigma_+^2} & \text{if } \mu \geq \beta, \\ \frac{(\mu-\beta)^2}{2\sigma_-^2} & \text{if } \mu < \beta. \end{cases} \quad (6.7)$$

Here, \mathcal{A} is the logarithm of the prior's normalisation constant, and cancels out in the detailed balance condition. For parameters with symmetric constraints, $\sigma_+ = \sigma_-$, then $\ln P(\mu)$ reduces to a quadratic with maximum at $\mu = \beta$.

6.1.7 Including Systematics in the Fit

One important implementation detail is how systematics are applied within the MCMC fitting process. Once systematics are added to the fit, at every step the binned PDFs for all the processes considered in the fit must get modified appropriately, which can become extremely computationally-intensive if not approached carefully. The strategy used in the `OXO` framework starts by thinking of the contents of a binned PDF as a vector of bin probabilities, $\mathbf{p} = (p_1, p_2, \dots, p_{N_{\text{bins}}})^T$. Then, we can think of a systematic acting on the PDF as a linear transformation, and hence a matrix S acting

Fig. 6.10

on this vector: $\mathbf{p}' = S\mathbf{p}$. We only need to calculate this matrix once for a given set of systematic parameter values, and can then use the same matrix on all the PDFs in the fit. Furthermore, when multiple systematics are applied, the matrix for each systematic can then be combined via matrix multiplication into one single “detector response” matrix. OXO uses the `Armadillo` [] linear algebra package for efficient matrix manipulation.

There is a problem that can arise when considering the impact of systematics near the edge of the analysis ROI. Many systematics such as shifts, scalings, and convolutions use information about the contents of nearby bins to determine the contents of a particular bin. However, for bins near the edge some of that information does not exist — it has been lost to the cuts that define the ROI. This can lead to a bias in the generation of the modified PDFs, and therefore also the posterior distribution.

As an example, consider the impact of an energy scale systematic on the energy distribution of ^{234m}Pa events in the detector, shown in Fig. 6.10. Because the events seen for this process in the ROI are merely the high-energy tail, any systematic energy scaling $E'_{\text{reco}} := \beta E_{\text{reco}}$ should have a large impact on the number of events observed at the low end of the ROI. However, given that the information about data below E_{min} is lost to the ROI cuts, any energy scaling of $\beta > 1$ will not be applied correctly at all.

The solution to this problem is defining a “buffer region” of bins on either side of the ROI, which allow for tracking of events in and out of the ROI due to systematics, but aren’t considered when calculating the likelihood. This is also shown in Fig. 6.10. After the scaling systematic is applied, although incorrect bin values are found in the buffer region, this is fine because we are no longer calculating the likelihood with those bins. Note that because of this modification, the normalisation parameters we put into the model no longer represent the expected number of events in the ROI. Instead, they

represent the number of events expected in both the ROI and buffer region, before any systematics have been applied.

6.1.8 Including Oscillations in the Fit

Within the analysis MCMC code, the process of neutrino oscillations are thought of as a *de facto* systematic that acts only on the ${}^8\text{B}$ ν_e and ν_x signal spectra. Three parameters relevant to the signal are floated within the MCMC fit: Δm_{21}^2 , θ_{12} , and $\Phi_{8\text{B}}$, the unoscillated ${}^8\text{B}$ flux relative to the expected rate. For the two signal PDFs, a third “bookkeeping” dimension is added on top of reconstructed energy and radius: the true neutrino energy, E_ν . This is necessary for correctly applying oscillations, as the oscillation probability is a function of the neutrino’s energy, not the scattered electron’s. Before the fit, these 3D PDFs are given normalisations corresponding to the expectation of the number of events for each type, ν_e and ν_x , after cuts but before oscillations have been applied. Strictly speaking there should be zero ν_x events before neutrino oscillations: the pre-oscillation rate used here is the post-cut number of events expected if 100% of the neutrinos oscillated to the ν_x type.

During the MCMC fit, for a given set of parameters $\boldsymbol{\theta} = (\Delta m_{21}^2, \theta_{12}, \Phi_{8\text{B}})$ the following is performed to oscillate the signal PDFs. Firstly, the normalisations are scaled by the factor $\Phi_{8\text{B}}$. Then, for each E_ν bin the survival probability $P_{ee}(E_\nu, \Delta m_{21}^2, \theta_{12})$ is calculated. Each bin then has their probability scaled by either P_{ee} or $1 - P_{ee}$, for ν_e and ν_x respectively. Of course, within the structure of the **OXO** framework these bin-by-bin scaling aren’t immediately applied, but instead a matrix describing the impact of oscillations on each of the PDFs is made. Because the oscillation transformation is purely a bin-by-bin scaling, the resulting matrices are diagonal, with diagonal elements $\Phi_{8\text{B}} \cdot P_{ee}(E_\nu, \Delta m_{21}^2, \theta_{12})$ or $\Phi_{8\text{B}} \cdot (1 - P_{ee}(E_\nu, \Delta m_{21}^2, \theta_{12}))$ for ν_e and ν_x respectively. After the oscillation matrix along with all other systematic matrices are applied to the

¹ signal PDFs, the PDFs are then marginalised over the E_ν dimension so that the signal
² PDFs match the dimensionality of all other processes.

³ Calculations of the survival probability are handled with **PSelmaa**, an algorithm
⁴ written by Nuno Barros for the SNO 3-phase Analysis []. This considers not only the
⁵ neutrino oscillations through the vacuum of space between the Sun and Earth, but also
⁶ the impact of matter effects in both the Sun and Earth. This can usually be a very
⁷ computationally-intensive process, but **PSelmaa** takes advantage of the assumption
⁸ that the solar oscillation parameters are in the so-called “Large Mixing Angle” regime,
⁹ making the calculation much faster. As seen in Section 1.2.1, previous solar oscillation
¹⁰ experiments demonstrate that this assumption is reasonable. For this analysis, the
¹¹ standard MSW effect is assumed with neutrinos obeying the Normal Hierarchy, with
¹² the Sun following the **B16_GS98** metallicity model [] and the **PREM** model being used
¹³ for the Earth [].

¹⁴ One final thing **PSelmaa** needs to know to calculate survival probabilities is the
¹⁵ distribution of solar zenith angles during the data-taking. The solar zenith θ_z is the
¹⁶ angle between the two following vectors: one going from the centre of the Earth through
¹⁷ the centre of the SNO+ detector, and another starting from the detector’s centre and
¹⁸ pointing towards the Sun. As an example, if the Sun were ever to be directly overhead
¹⁹ the detector, both vectors would be along direction \hat{z} in detector coordinates, leading
²⁰ to a solar zenith angle of $\theta_z = 0$. The position of the SNO+ detector on Earth, as
²¹ well as the times at which the detector was live, determine the solar zenith angle
²² distribution. If not accounted for, this can lead to a bias in the result of the analysis,
²³ as a preponderance of livetime taken at night (say) would lead to a larger fraction of
²⁴ solar neutrinos having to pass through the bulk of the Earth to get to the detector,
²⁵ and hence the impact of the MSW effect would be greater.

6.2 Analysis on Scintillator-Phase data

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Even after using the Large Mixing Angle approximation, having to call `PSelmaa` numerous times for every step in the MCMC algorithm would lead to exorbitant run times for the fitting. Therefore, a further approximation is made. Before running the MCMC fit, `PSelmaa` is used to calculate P_{ee} over the necessary 3D space of parameters. To get a fine scan of this space, 101 E_ν values from 1 MeV to 20 MeV, 101 Δm_{21}^2 values from 3×10^{-6} eV² to 1×10^{-3} eV², and 151 values for θ_{12} from 5° to 65° were looked over. This 3D grid of $101 \cdot 101 \cdot 151$ P_{ee} values is then written to disk, and loaded into memory for use during the fit as a lookup table. At run-time, as the Metropolis-Hastings algorithm samples this 3D space the survival probability is estimated through a trilinear interpolation of the 3D grid loaded in: a version of linear interpolation for three dimensions [].

[19 pages currently, without figures; probably +5 pages for figures.]

6.2 Analysis on Scintillator-Phase data

6.2.1 Dataset and Livetime

- Description of dataset chosen for analysis: 2.2 g/L scintillator-phase data that satisfies the “gold” list of run selection requirements, between May and November 2022.
- Explain requirements for run being selected into the Gold list.
- Note ‘raw’ livetime calculated for this dataset, and then calculate the impact that the muon and high-nhit vetos have on the livetime.
- Note which RAT versions MC production is being used to compare to data.

[2 pages]

6.2.2 Event Selection

- List final set of cuts chosen for analysis, along with any explanations for cuts that haven't already been motivated earlier (e.g. data cleaning). These are:
 - Event index cut (prevents MC events that don't trigger detector from being used)
 - Data cleaning cuts
 - High-nhit event timing veto cut
 - Valid scintFit reconstruction
 - Reconstructed energy $2 < E < 14$
 - Radial fiducial volume cut
 - BiPo out-of-window tag
 - BiPo in-window classifier cut
 - Externals classifier cut
 - “Cleanliness” cut
 - Position FOM cut
 - Show impact of cuts on data and MC. Show tables (the full details maybe in an appendix) indicating this.
 - Describe the finalised choice of binning for PDFs and data.
- [4 pages]

6.2.3 Expected Rates and Constraints

- Show expected rates calculation for both signal and background.

6.2 Analysis on Scintillator-Phase data

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- Describe the constraints chosen to apply to the fit, and why they can be justified. 1

These are:

- B8 flux constraint 3
- U-238 and Th-232 constraints from BiPo tagging 4
- Alpha-n constraint from Po-210 tagging by Serena and Shengzhao 5
- External constraints from Tony Zummico's water-phase externals analysis. 6
- Describe the systematics to be added to the fit (just energy scale for now, maybe also energy smearing?). For other possible systematics, such as those in position, my aim is to explain why they are sub-dominant and so don't need to be added to the fit. Will cover more about impact of systematics in the next subsection. 7
8
9
10

[4 pages]

11

6.2.4 Results

12

Fit Validation

13

- Show plots of parameter values versus step, to demonstrate that the step sizes have been tuned sufficiently. 14
15
- Show auto-correlation plots, to motivate a sensible “burn-in” size. 16
- Posterior density plots for each nuisance parameter, to check that they all look sensible and have sufficient statistics. 17
18
- Show plot of correlation coefficients between parameters, and note any strongly-correlated parameters. 19
20

[6 pages]

21

1 Oscillation Fit Results

- 2 • Look at the data versus ‘best-fit’ MC plot in energy, radius, and both. (Recall
3 that in MCMC, the ‘best-fit’ is not the point of parameter scape reached at the
4 end, but the point of highest posterior density). Is there a good fit to data? Any
5 clear disagreements?
- 6 • Show 2D contour plot for oscillation parameter posterior density. Note salient
7 features. Show 1D posterior densities for each oscillation parameter. Derive mea-
8 surement result for θ_{12} , i.e. point of highest posterior density, with uncertainties
9 given by the 1σ credible interval.

10 [5 pages]

11 Impact of Systematics

- 12 • Show impact of modifying certain constraints on the final results of the measure-
13 ment of θ_{12} . In particular: fiducial volume, ${}^8\text{B}$ flux constraint.
- 14 • Discussion of systematics post-fit — Hopefully energy scaling parameter should
15 be close to 1, given the Collaboration’s calibration of the optics (light yield in
16 particular). Perform a scan over energy smearing, and check impact of possible
17 radial scale systematic.

18 [8 pages]

19 6.3 Sensitivity Projections

- 20 • Using the same production MC, generate PDFs with the expected normalisations
21 for longer periods of livetime: 1, 3, and 5 years. This still assumes a scintillator-fill

with identical detector conditions on average, so not considering the impact of BisMSB loading at any point.

- Describe any further constraints to be assumed on top of the existing analysis of data: expected improved constraints on various backgrounds, as well as a possible energy-scale calibration constraint from the AmBe source or an internally-deployed source, for example.
- Run MCMC fits to these Asimov PDFs for each livetime scenario. Describe results in terms of the improvement to the sensitivity to θ_{12} as a function of livetime.
- Also consider scenario of lower backgrounds! How does that impact the result?
- Could consider scenario of BisMSB deployment, leading to substantially greater light yield and hence energy resolution. Given time it would take for Production MC to be generated for this scenario, I would likely have to come up with something clever and quick to actually do this. There quite possibly won't be enough time.

[8 pages]

[61 PAGES TOTAL]

Chapter 7

Conclusions and Suggestions for Future Work

- Say what has been achieved in this thesis! In particular:
 - Substantial improvement to the SMELLIE generator in terms of speed and dynamic range
 - A much stronger understanding of the discrepancies between data and MC in SMELLIE
 - The creation of two analyses of SMELLIE data, designed explicitly around being robust to these systematics
 - A measurement of the extinction length of scintillator *in-situ* with SMELLIE at 375 nm, monitored over time
 - A first measurement of the scattering length of the scintillator *in-situ*, monitored over time
 - The creation of an analysis of ${}^8\text{B}$ solar neutrinos in the scintillator phase to measure the solar neutrino oscillation parameters

- ¹ – The first measurement of θ_{12} using ^8B neutrinos in SNO+
- ² – Projections of this solar analysis' precision at longer livetimes
- ³ • Give suggestions for further work that could be done on both SMELLIE and the
- ⁴ solar oscillation analysis:
 - ⁵ – Inclusion of LABPPO's polarisability anisotropy in the detector's optical
 - ⁶ scattering simulation, and determination of its impact on both Physics and
 - ⁷ SMELLIE measurements
 - ⁸ – Further investigation of SMELLIE's wavelength-dependence of the beam
 - ⁹ profiles, looking both at possible origins for the phenomenon and correcting
 - ¹⁰ for this in simulation
 - ¹¹ – Various computational improvements to allow for faster MCMC run-times,
 - ¹² including the calculation of the systematic matrices, and possible use of
 - ¹³ GPUs to parallelise some parts of the computation.
 - ¹⁴ – Inclusion of additional solar neutrino components at lower energy into the
 - ¹⁵ fit, e.g. Be7. Maybe the addition of lower energies also helps to naturally
 - ¹⁶ constrain backgrounds within the fit?
 - ¹⁷ – Looking at the impact of various advanced background-rejection procedures,
 - ¹⁸ such as event directionality or topology. How much do they help with the
 - ¹⁹ sensitivity?
 - ²⁰ – Looking at the impact of splitting data into day and night parts, to try
 - ²¹ and provide further constraints on any matter effects. Not expected to be
 - ²² significant, so was ignored for this analysis so far.
 - ²³ – Performing a combined fit with the reactor anti-neutrino oscillation analysis,
 - ²⁴ which allows for the handling of correlated uncertainties, such as detector
 - ²⁵ response systematics.

[3 PAGES TOTAL]

1

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