

Development of Photon Calibrator for Hardware Injection Test

Yu-Kuang Chu
2018 Oct. 25

outline

- Introduction
- Hardware Injection through Photon Calibrator
- Signal Generating System (Digital Control System)
- Noise reduction through De-Whitening Filter
- Validation of Injection Channel
- Discussion and Future Works

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Development of Photon Calibrator for Hardware Injection Test

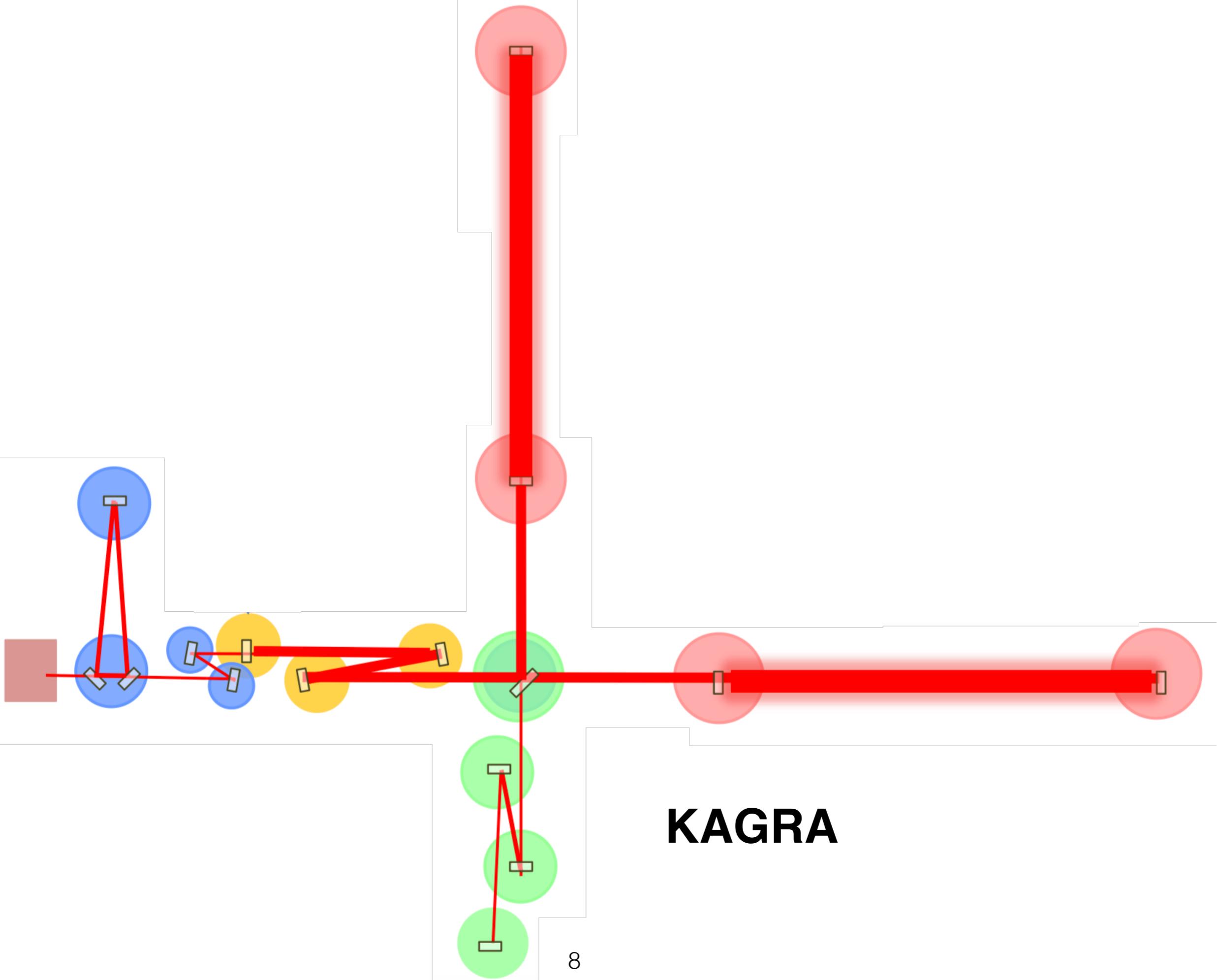
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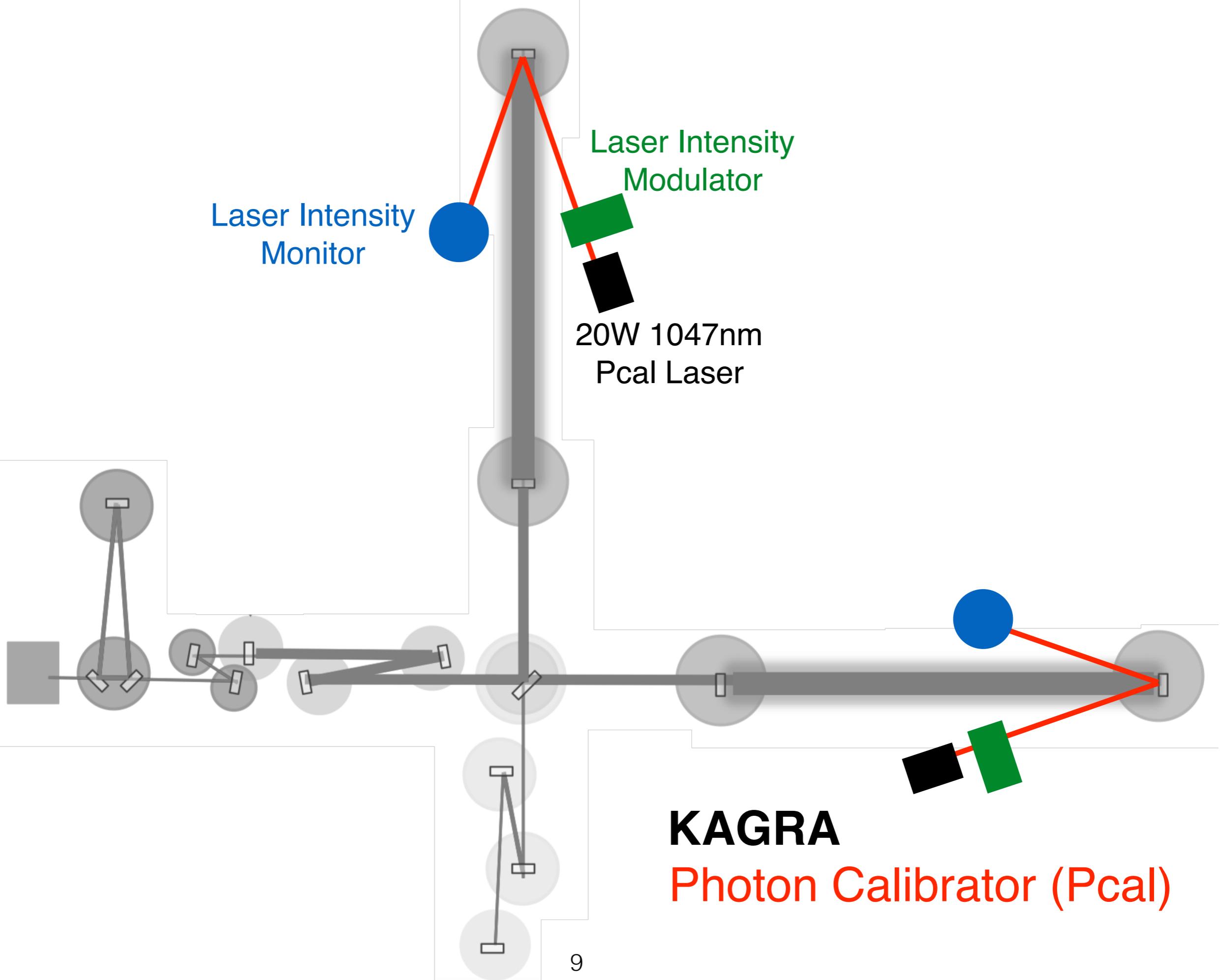
What is Hardware Injection Test?

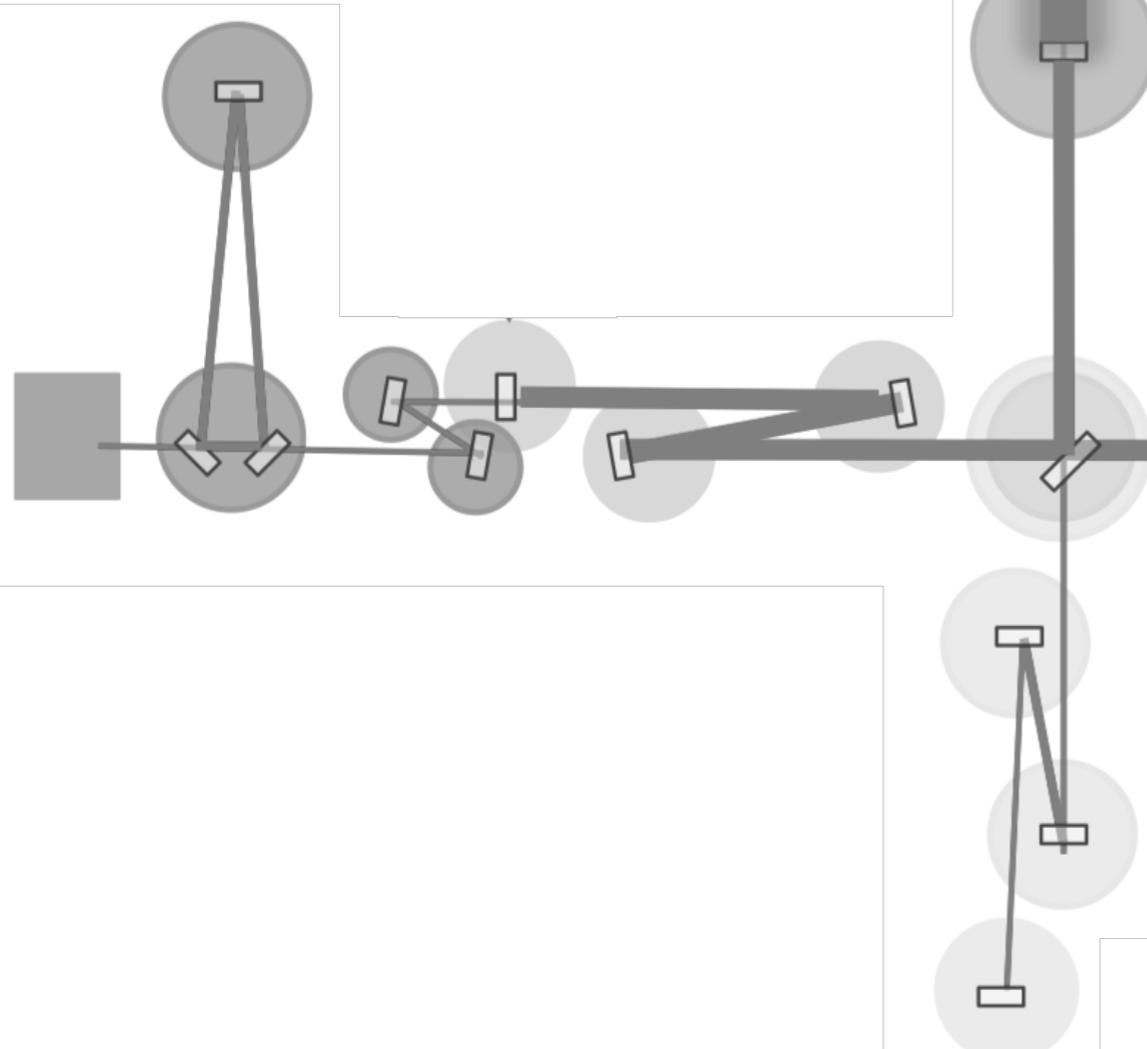
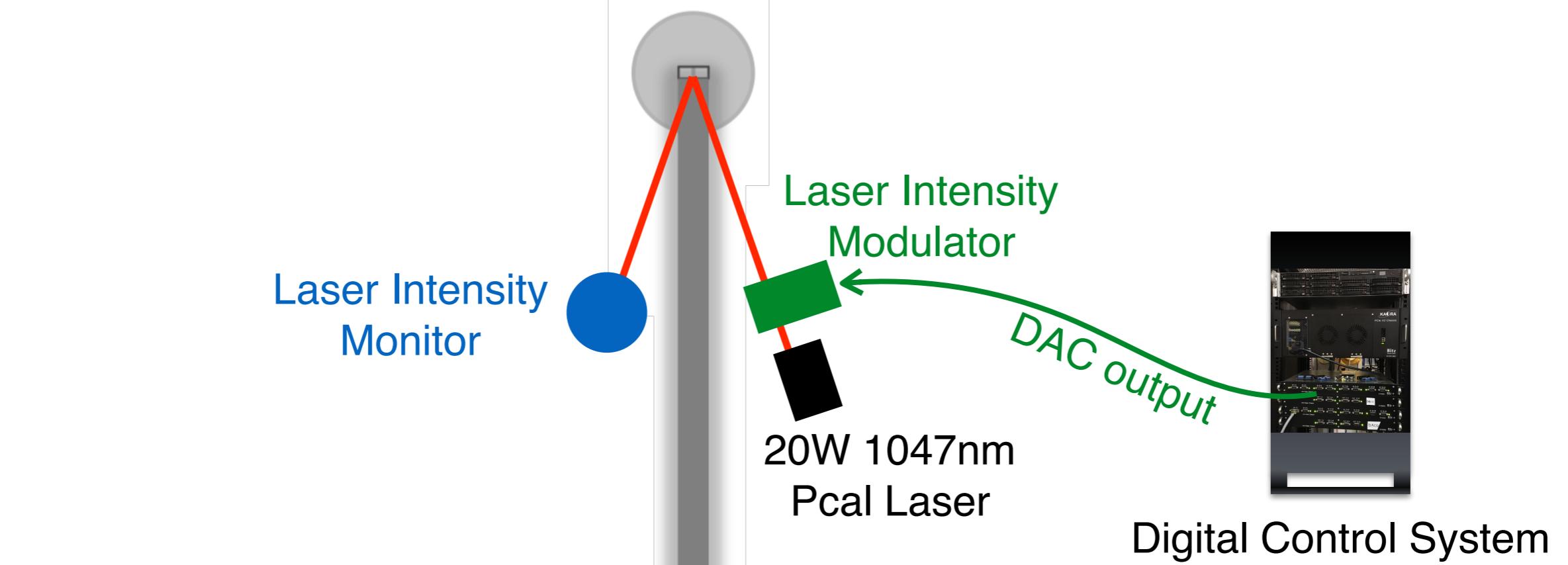
Hardware injection test is a **process to verify the performance of interferometer by sending sample signals into interferometer**. Ideally, we should prepare some real gravitational waves as test signals. But it is practically hard to generate large enough artificial gravitational waves that are detectable by current technology.

Therefore, instead of generating gravitational waves, people mimic the effect of celestial gravitational waves by **displacing the mirror according to the simulated gravitational waveforms**, changing arm length correspondingly, comparing the optical readout in the main interferometer, thereby checking the response of their interferometer.

What is Photon Calibrator(PCal)?







Calibration Lines

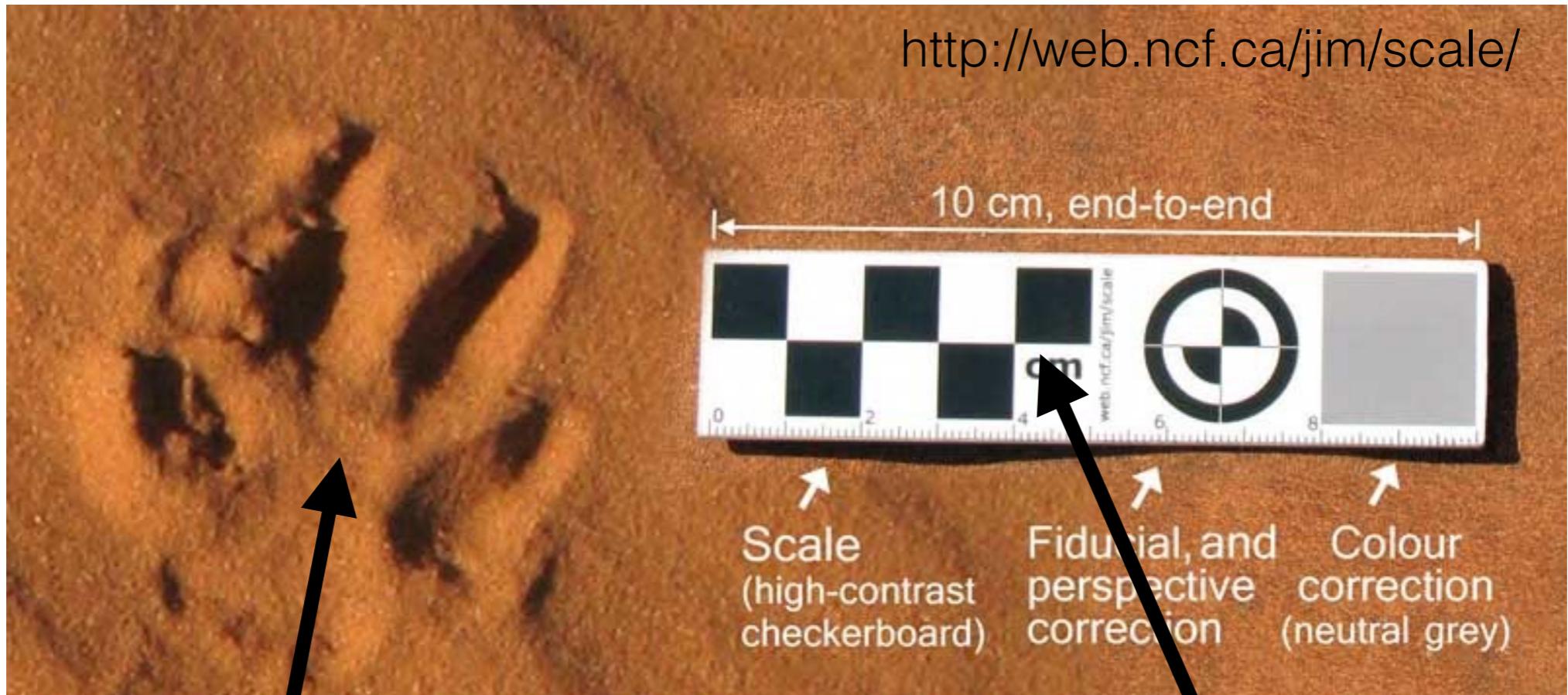
Single frequency sine wave as fiducial length displacement

Theoretical GW waveform

Binary Neutron Stars, Binary Blackholes

Swept-Sine Excitation

Transfer Function Measurement



Interferometer readout
generated by
gravitational wave signal
(GW signal)

Interferometer readout
generated by Mirror
displacement due to
Photon Calibrator
(Calibration Lines)

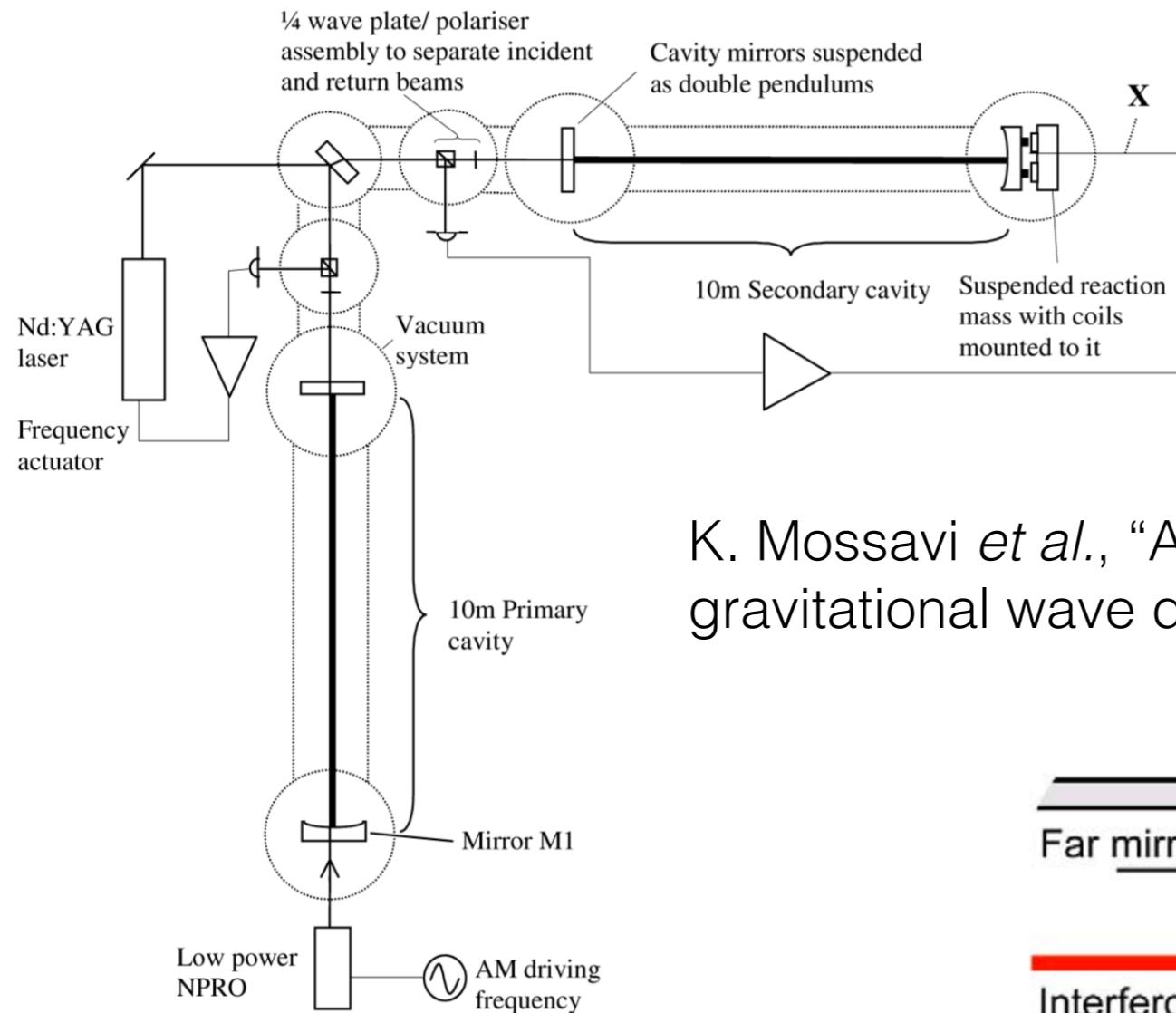
Interferometer readout
generated by
gravitational wave signal
(GW signal)



Interferometer readout
generated by Mirror
displacement due to
Photon Calibrator
(Calibration Lines)

History of Photon Calibrators

D. A. Clubley *et al.*, “Calibration of the **Glasgow 10 m prototype** laser interferometric gravitational wave detector using photon pressure,” Phys. Lett. A 283, 85 (2001).



K. Mossavi *et al.*, “A photon pressure calibrator for the **GEO600** gravitational wave detector,” Phys. Lett. A 353, 1 (2006).

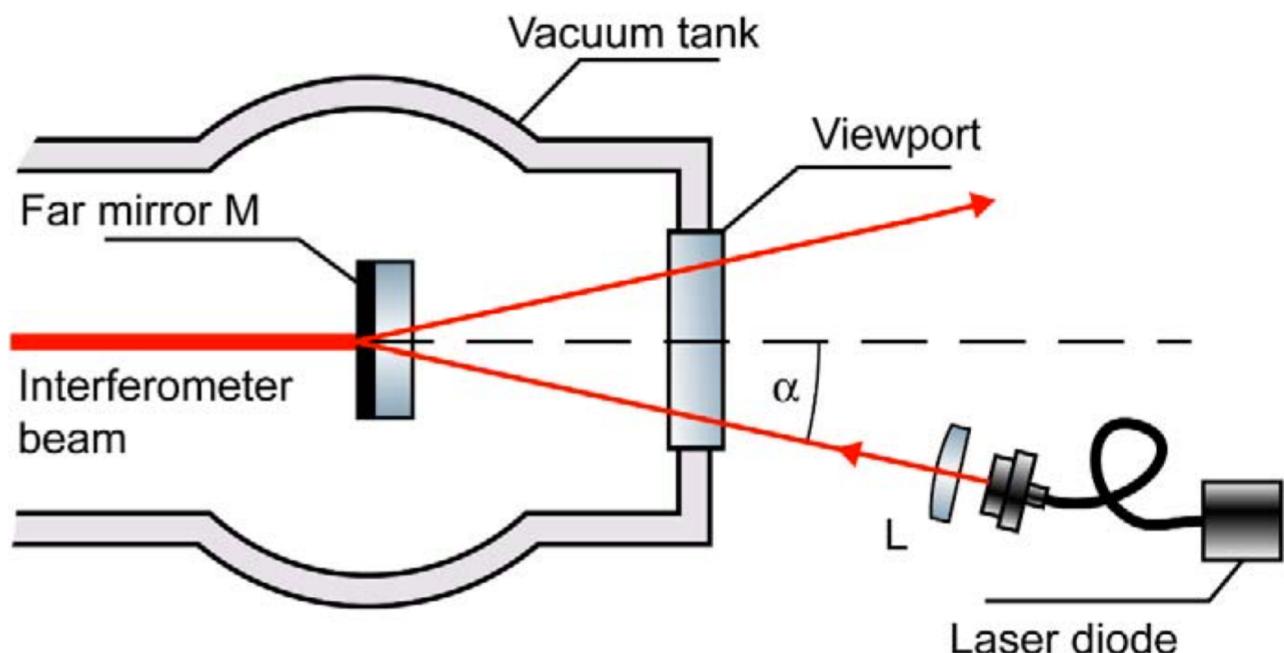


Fig. 1. Experimental arrangement for measurements with the photon calibrator. M: Far mirror in the north building of the GEO 600 detector. Illumination of this mirror with light from the laser diode produces a differential arm-length change that is measured at the main output of the interferometer.

S. Hild et al., “Photon pressure induced **test mass deformation** in gravitational wave detectors,” Classical Quantum Gravity 24, 5681–5688 (2007)

Badan M A, Landry M, Savage R and Willems P (2009) Analyzing **elastic deformation** of test masses in LIGO LIGO doc. T0900401

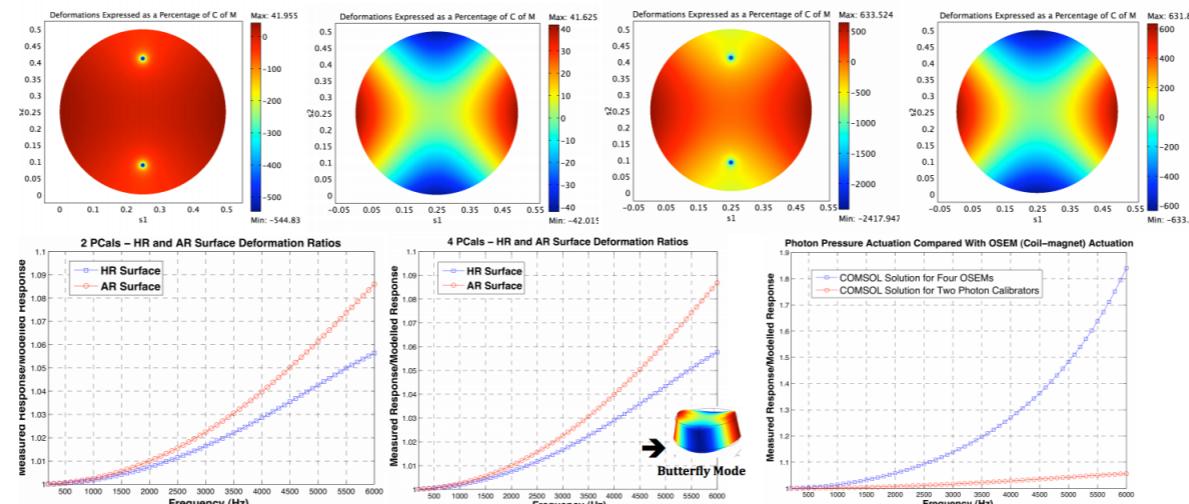
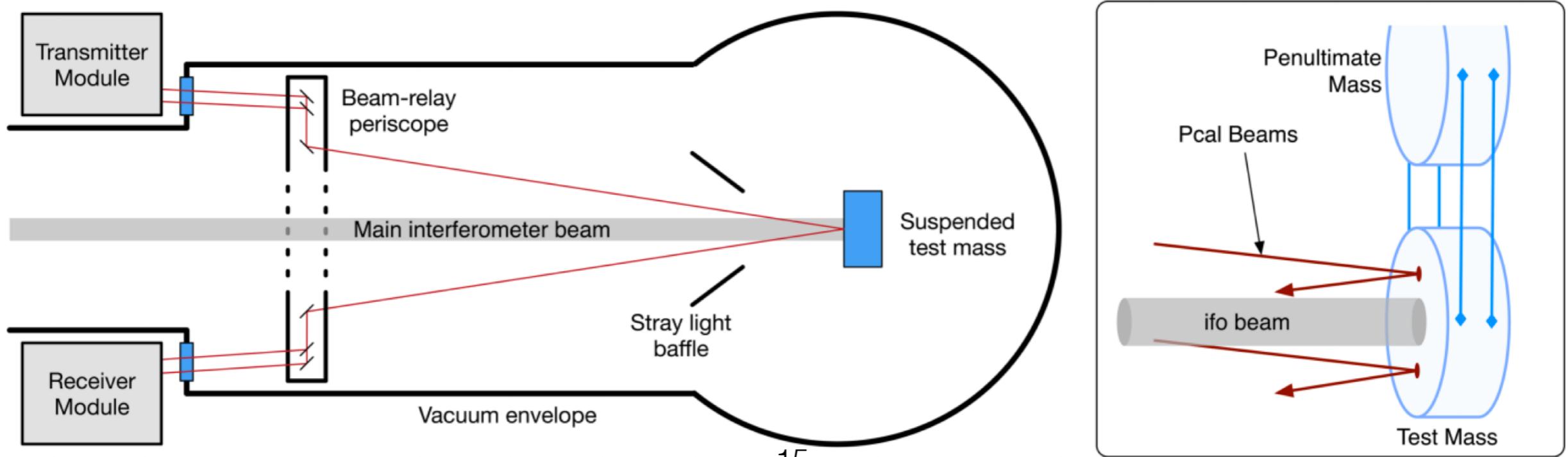
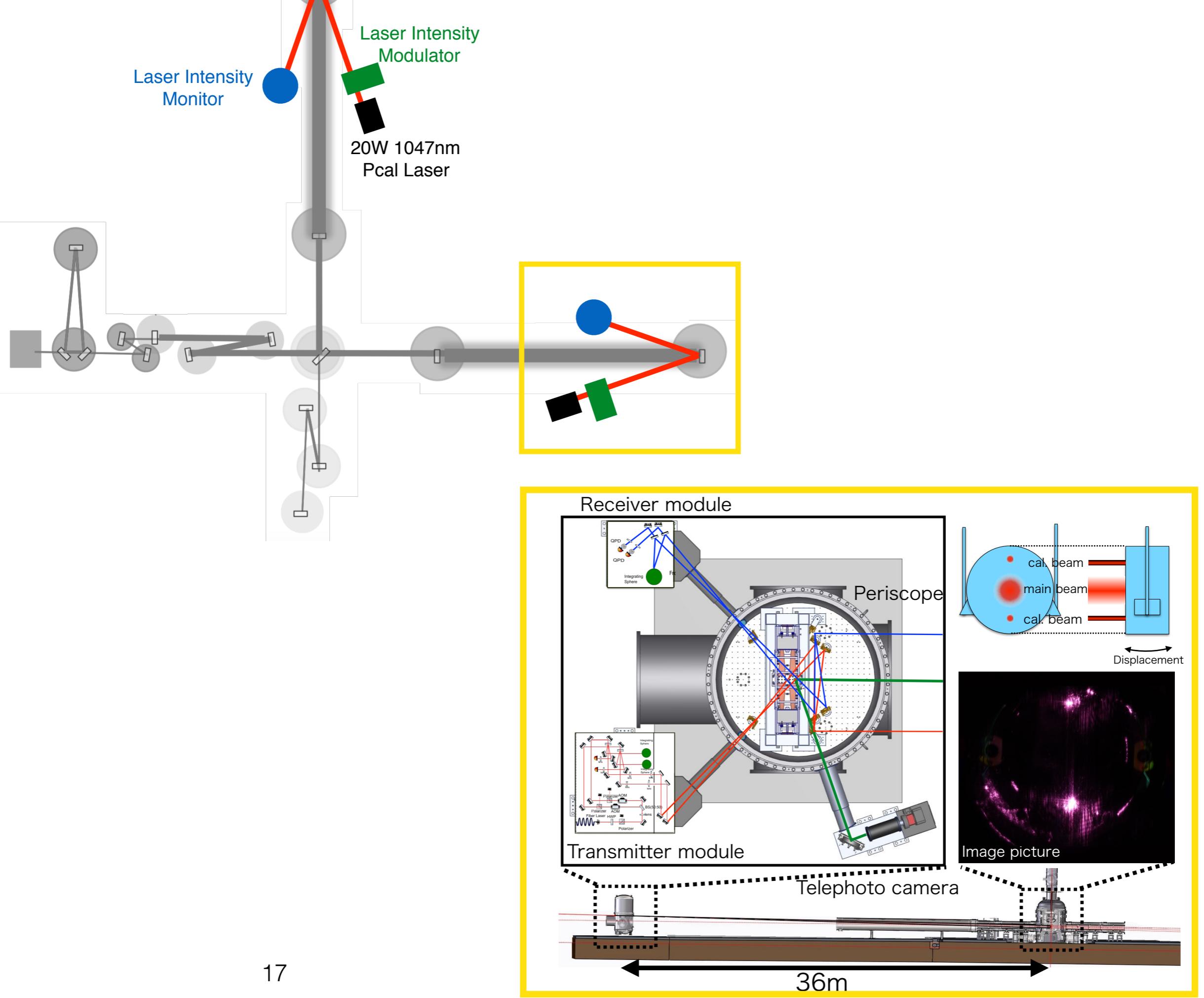


Figure 6 | (Top) The HR (left) and AR (right) z-displacements for the photon actuation. The graphs show the 3D surface deformation as a percentage of the center of mass motion for 3 kHz and 6 kHz, respectively. (Bottom) The HR and AR surface Measured/Modelled Ratio profiles for the two photon calibrator forces, followed by four. The discrepancies experienced are about are about an order of magnitude less than the OSEM case for both configurations. This is because the optic flexes differently in perpendicular directions as they try to form the butterfly mode at 6823 Hz.

S. Karki, et al., “**The Advanced LIGO Photon Calibrators**”, Rev. Sci. Instrum. 87, 114503 (2016).



KAGRA Photon Calibrator



Receiver module

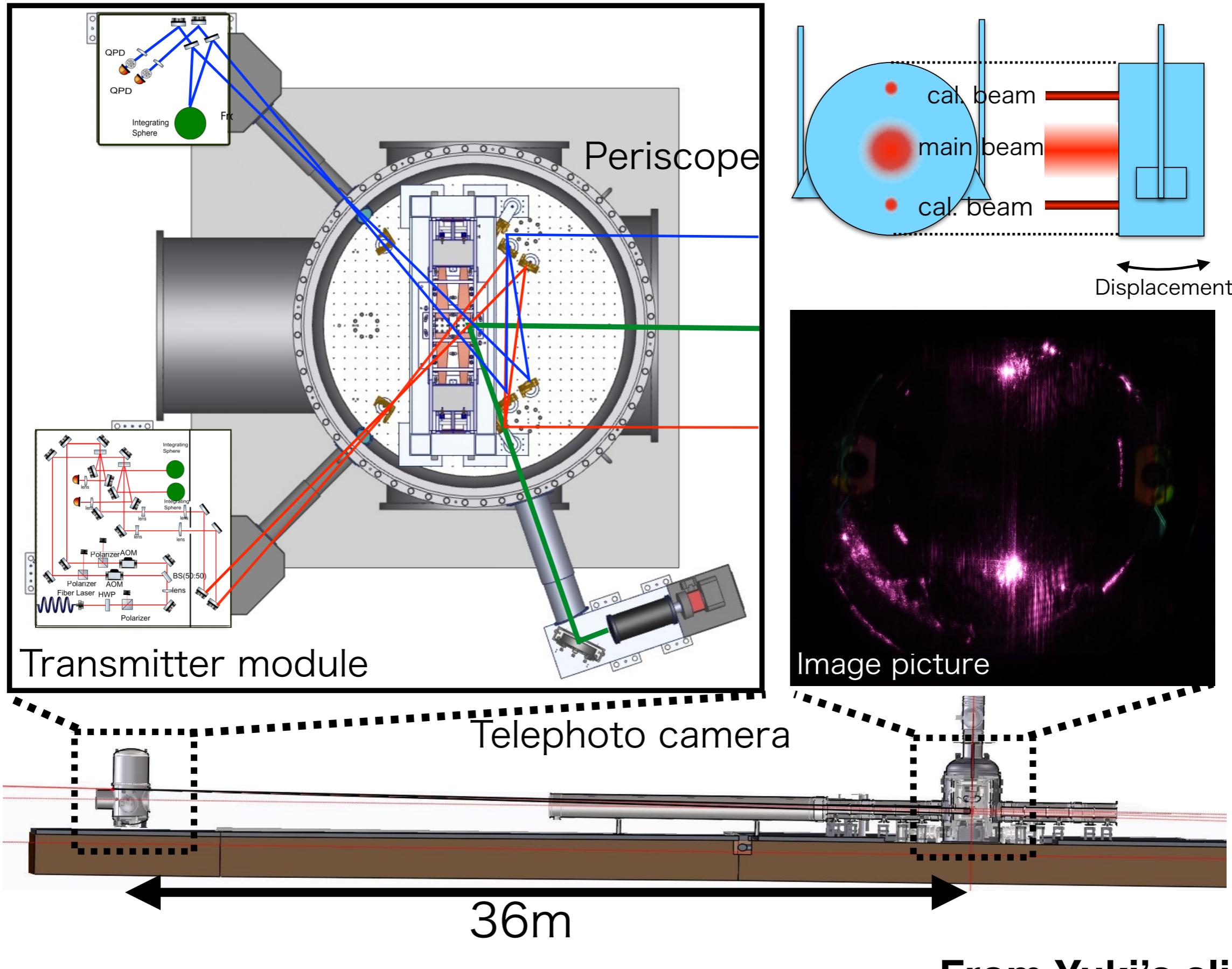
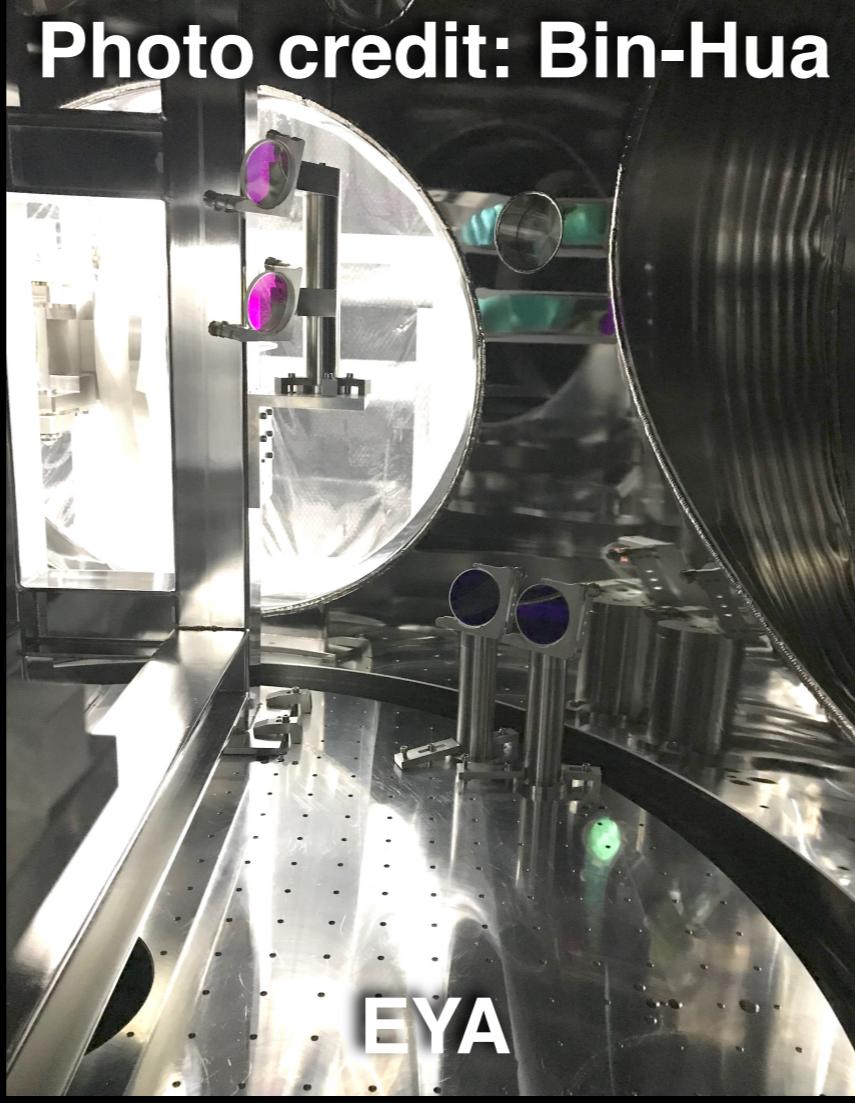


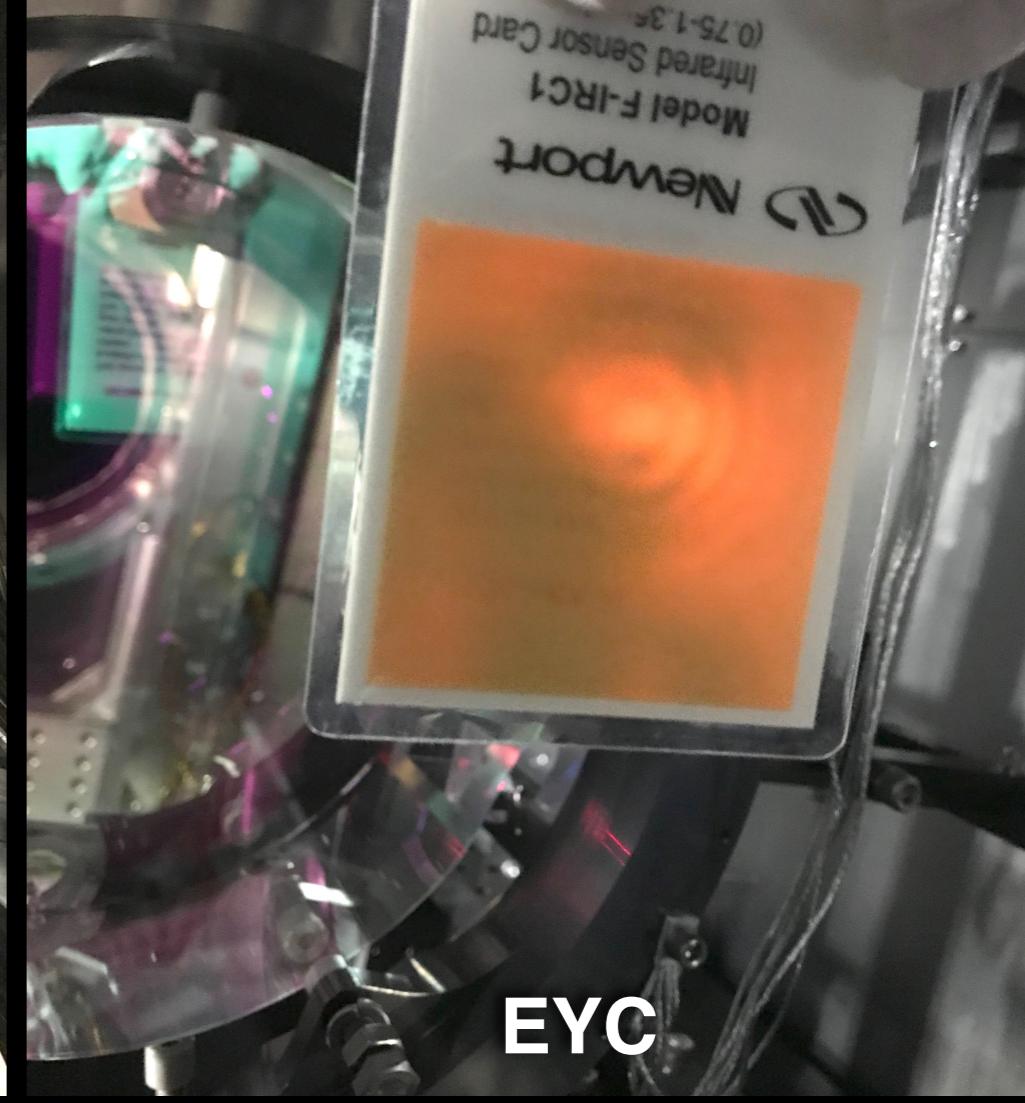
Photo credit: Bin-Hua



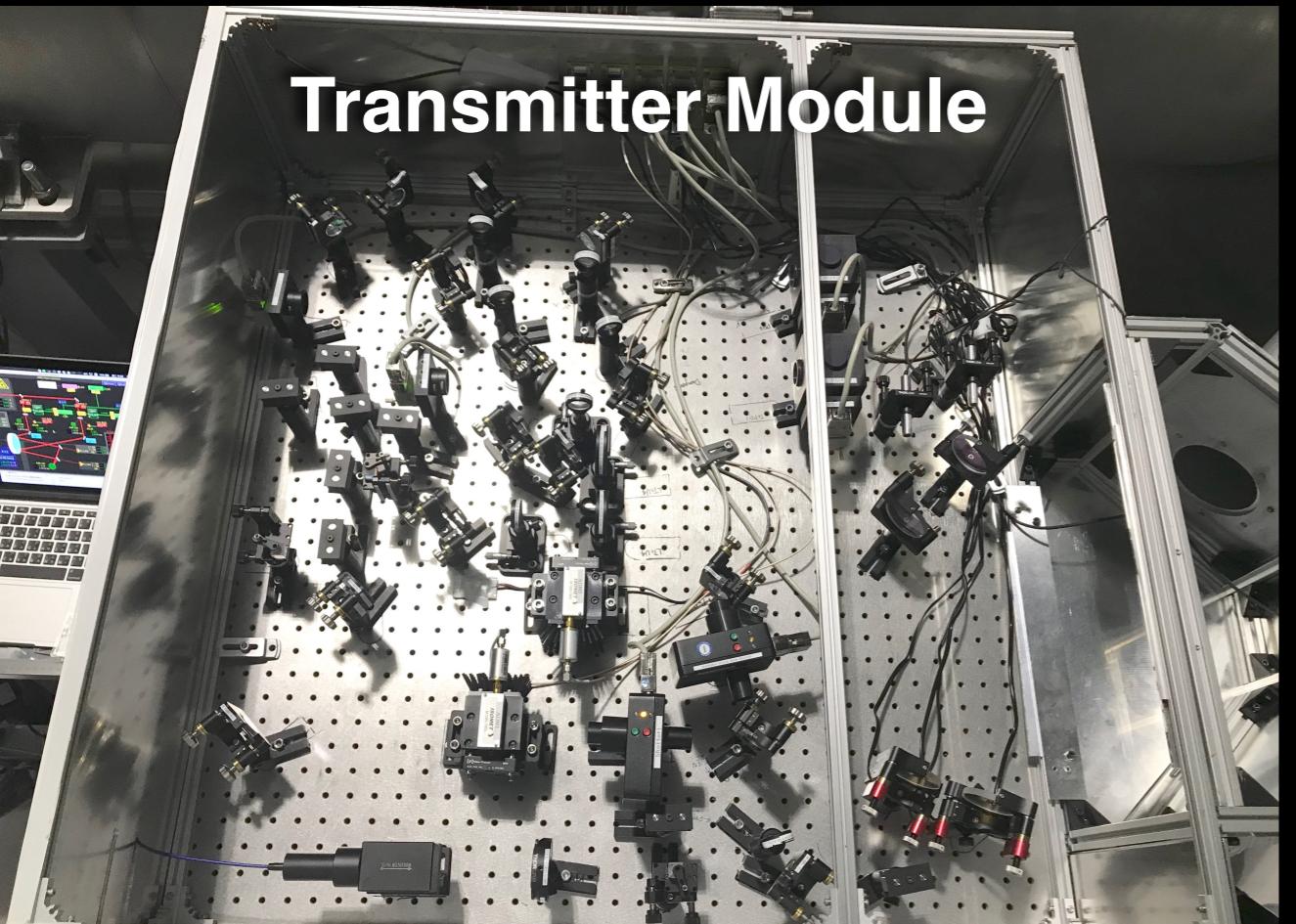
Receiver Module



EYA



EYC



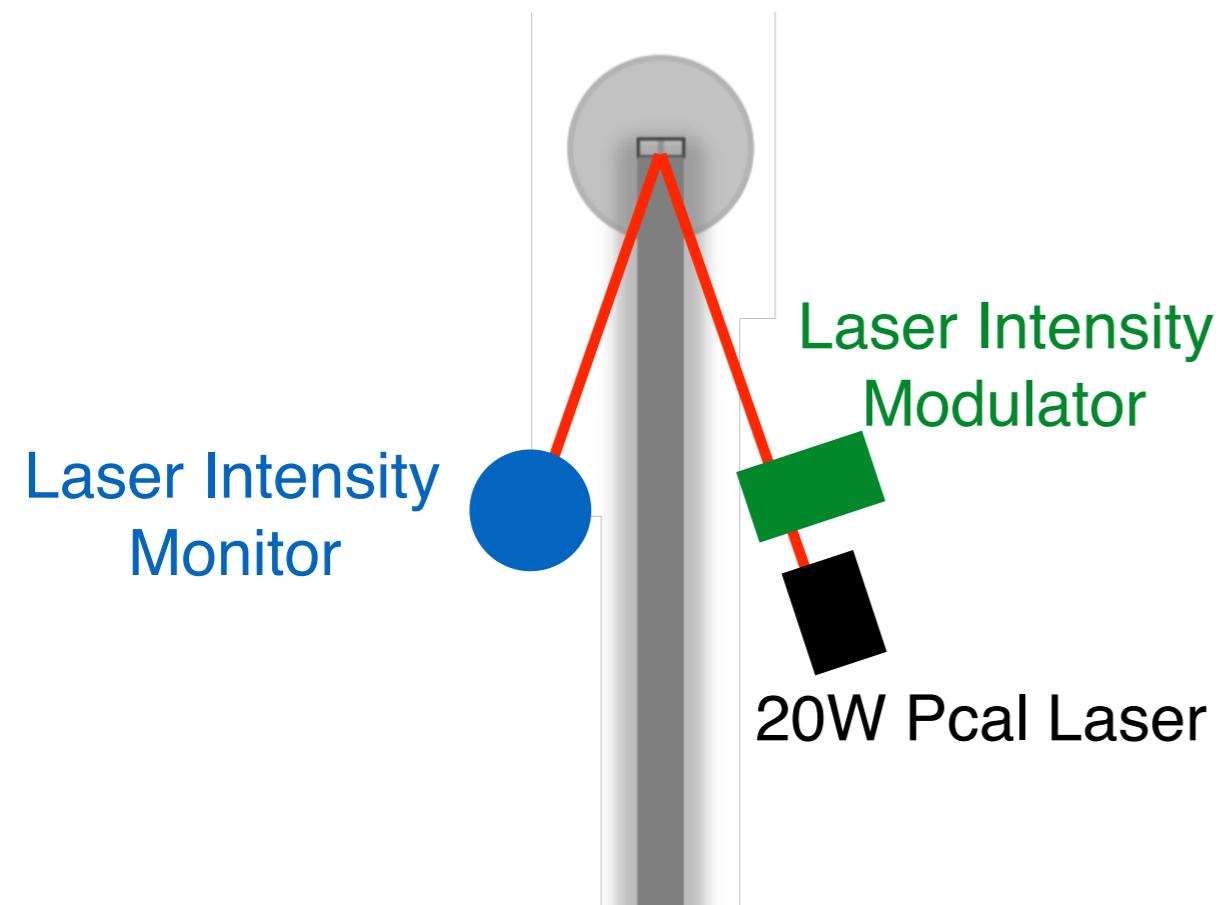
Transmitter Module



Transmitter Module

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$$\ddot{x}(t) + b\dot{x}(t) + \omega_0^2 x(t) = \frac{F(t)}{M}$$

$$x(\omega) = \frac{-1}{\omega^2 - \omega_0^2 - i\omega b} \frac{F(\omega)}{M}$$

$$\Delta L(f) = \underbrace{\frac{2\Delta P(f) \cos(\theta)}{c}}_{\text{Interferometer Mirror displacement induced by Pcal Laser intensity modulation}} \underbrace{\frac{-1}{M(2\pi f)^2}}_{\substack{\text{Radiation Force exerted by Pcal} \\ \text{Force to length transfer function}}}$$

Interferometer Mirror displacement
induced by Pcal Laser intensity
modulation

Radiation Force exerted by Pcal
Force to length transfer function

Theoretical waveforms

$$\int_{-\infty}^{-\infty} x(\omega) e^{i\omega t} d\omega = \frac{-2 \cos(\theta)}{Mc} \int_{-\infty}^{-\infty} \frac{P(\omega)}{\omega^2} e^{i\omega t} d\omega$$

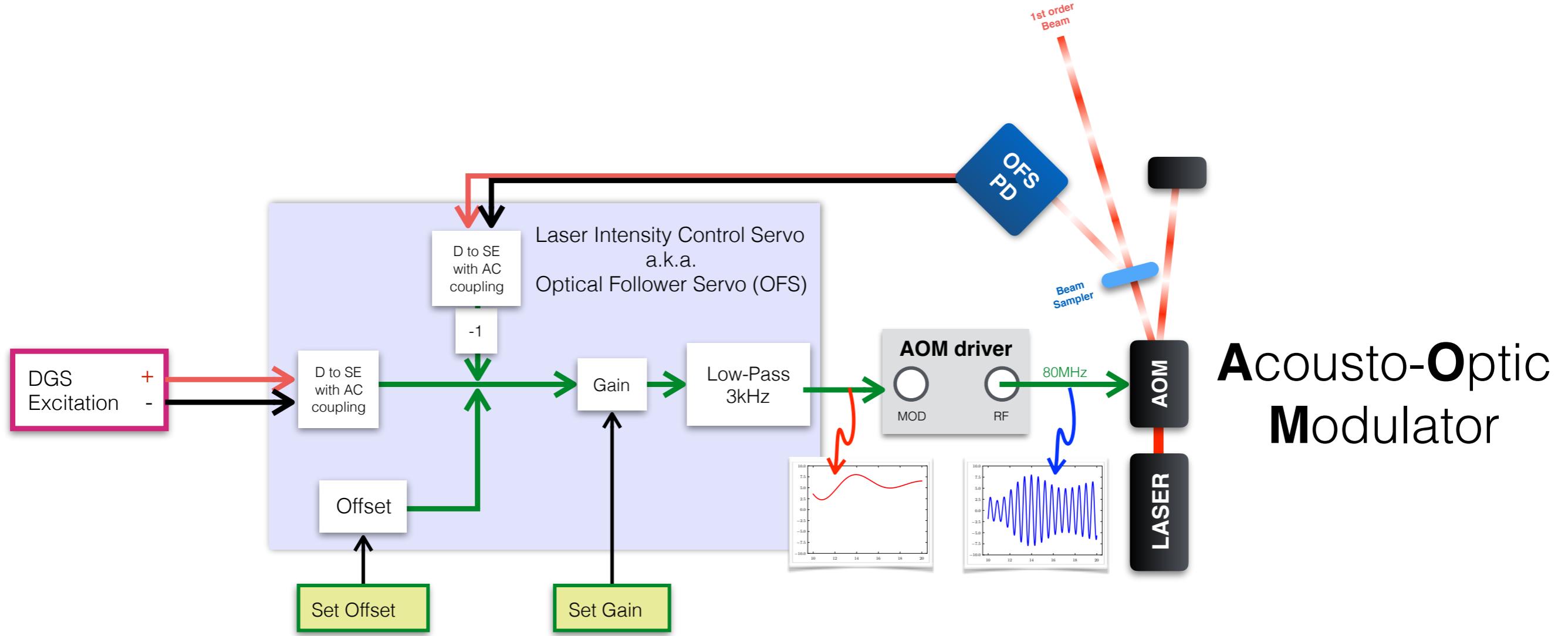
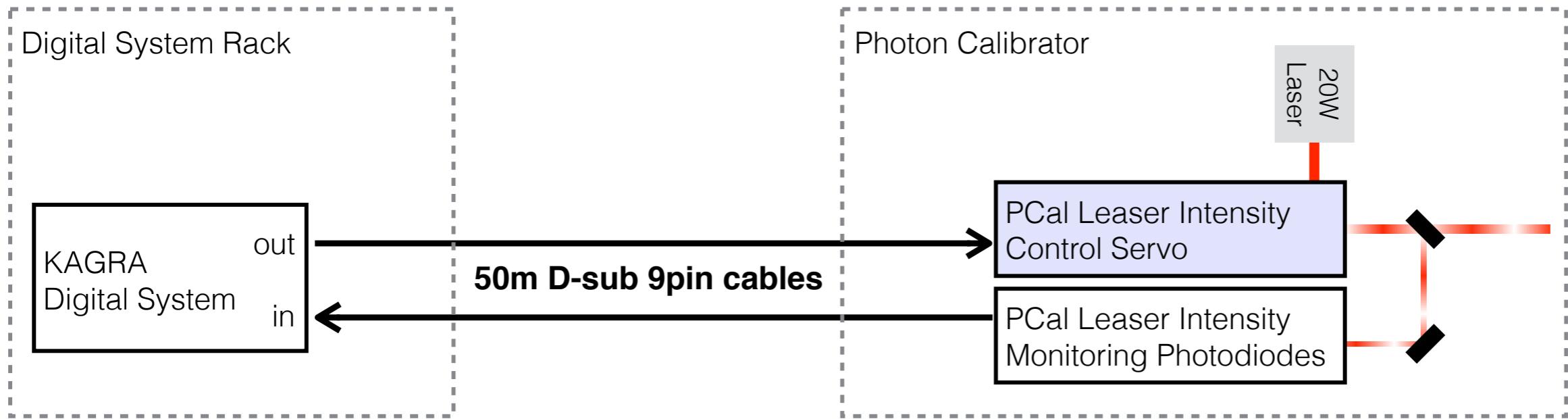
$$\frac{Mc}{-2 \cos(\theta)} \int_{-\infty}^{-\infty} x(\omega) \omega^2 e^{i\omega t} d\omega = P(t)$$

Required PCal laser intensity modulation

$$\Delta L(f) = \underbrace{\frac{2\Delta P(f) \cos(\theta)}{c}}_{\text{Interferometer Mirror displacement induced by Pcal Laser intensity modulation}} \underbrace{\frac{-1}{M(2\pi f)^2}}_{\substack{\text{Radiation Force exerted by Pcal} \\ \text{Force to length transfer function}}}$$

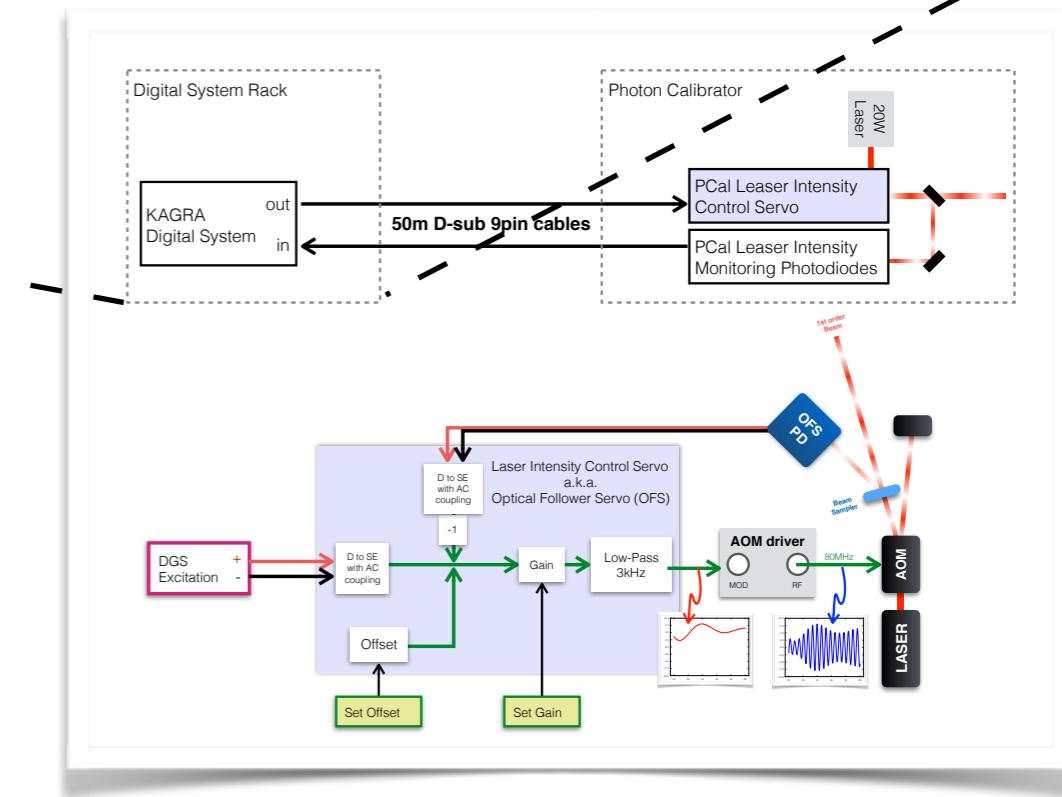
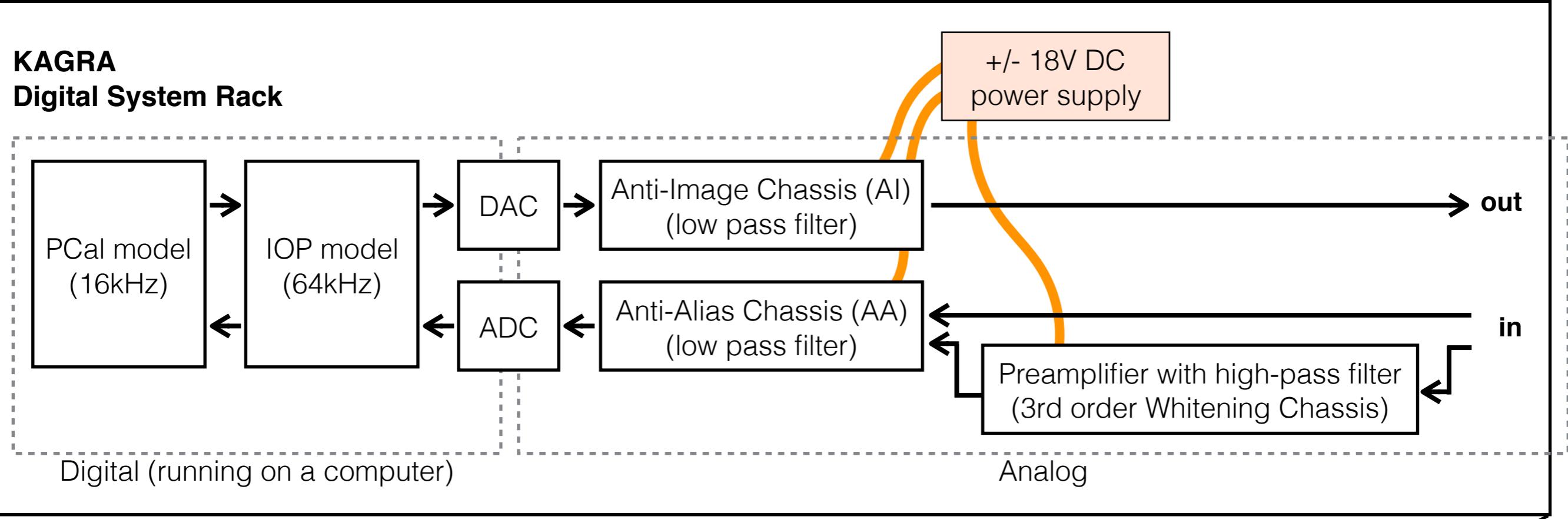
Interferometer Mirror displacement induced by Pcal Laser intensity modulation

Radiation Force exerted by Pcal Force to length transfer function



KAGRA

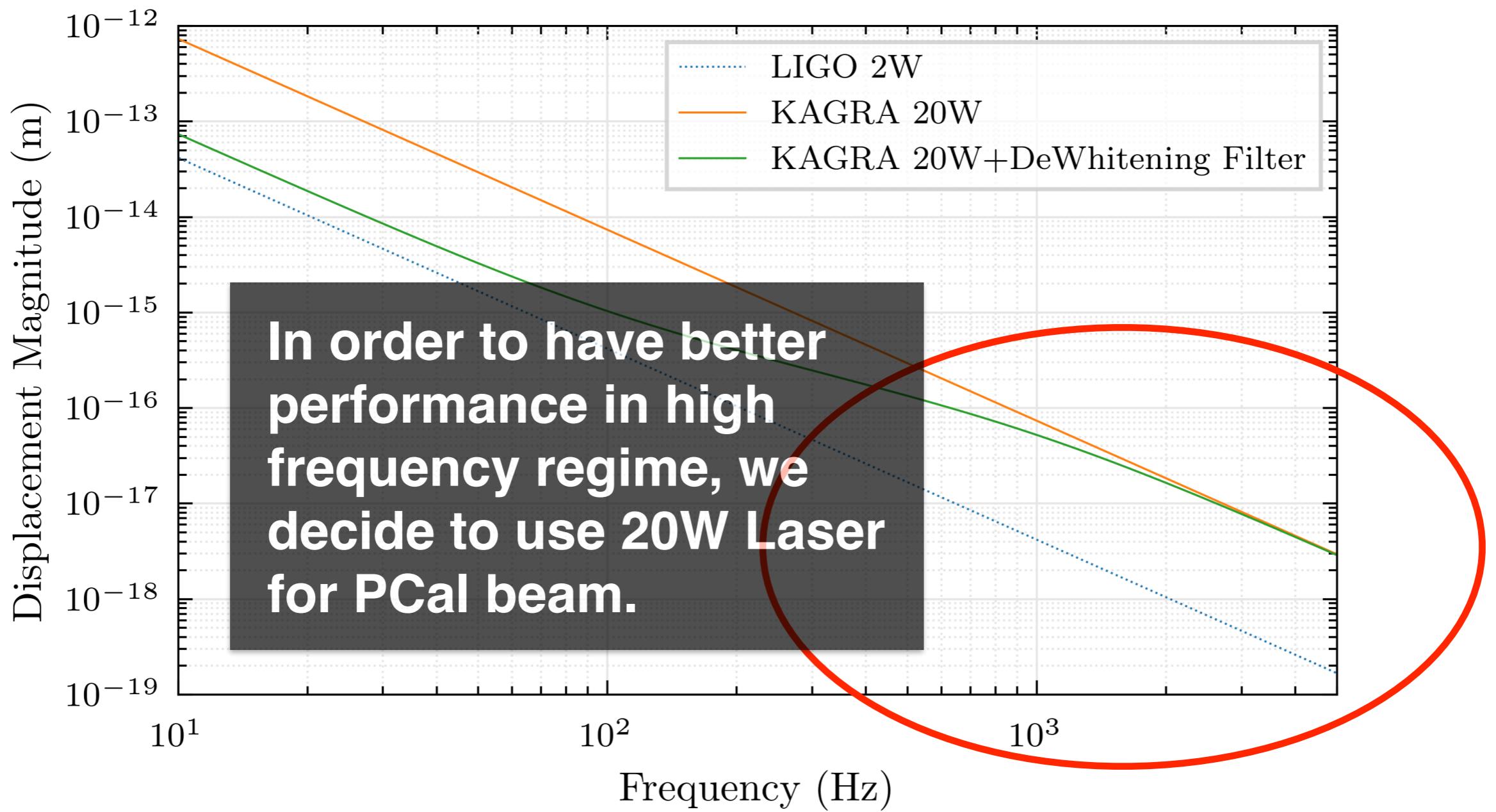
Digital System Rack



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Excitation Signal Noise Problem



$$\Delta L(f) = \underbrace{\frac{2\Delta P(f) \cos(\theta)}{c}}_{\text{Radiation Force exerted by Pcal}} \underbrace{\frac{1}{M(2\pi f)^2}}_{\text{Force to length transfer function}}$$

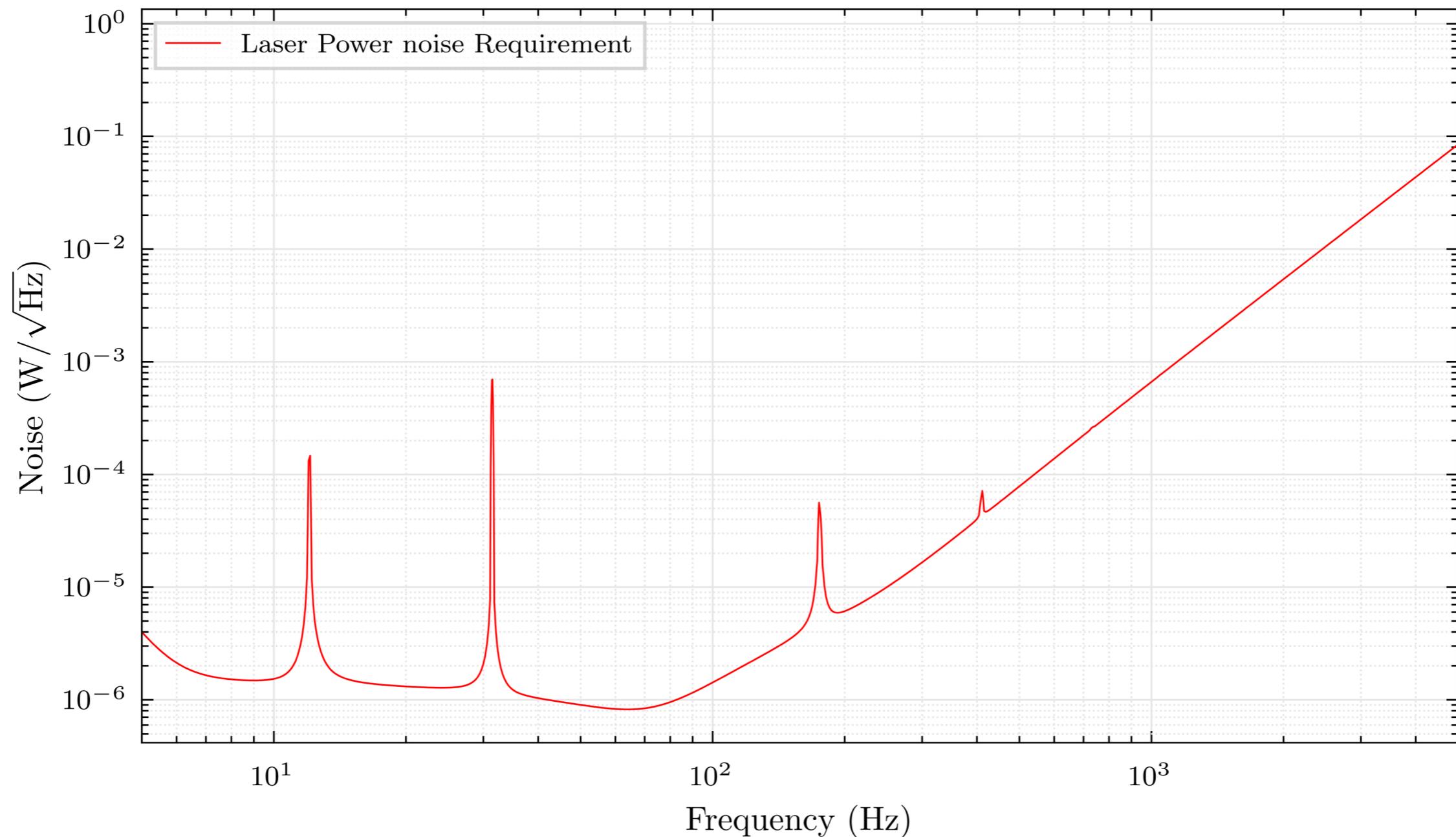
Interferometer Mirror displacement induced by Pcal Laser intensity modulation

Radiation Force exerted by Pcal

M = KAGRA 23kg
aLIGO 40kg

θ = KAGRA 0.72 deg
aLIGO 8.75 deg

Laser Power Noise Requirement



$$\Delta L(f) = \frac{2\Delta P(f) \cos(\theta)}{c} \frac{1}{M(2\pi f)^2} < \frac{1}{10} \Delta h(f)L$$

$\Delta L(f)$

Interferometer Mirror displacement induced by Pcal Laser intensity modulation

$\Delta P(f) \cos(\theta)$

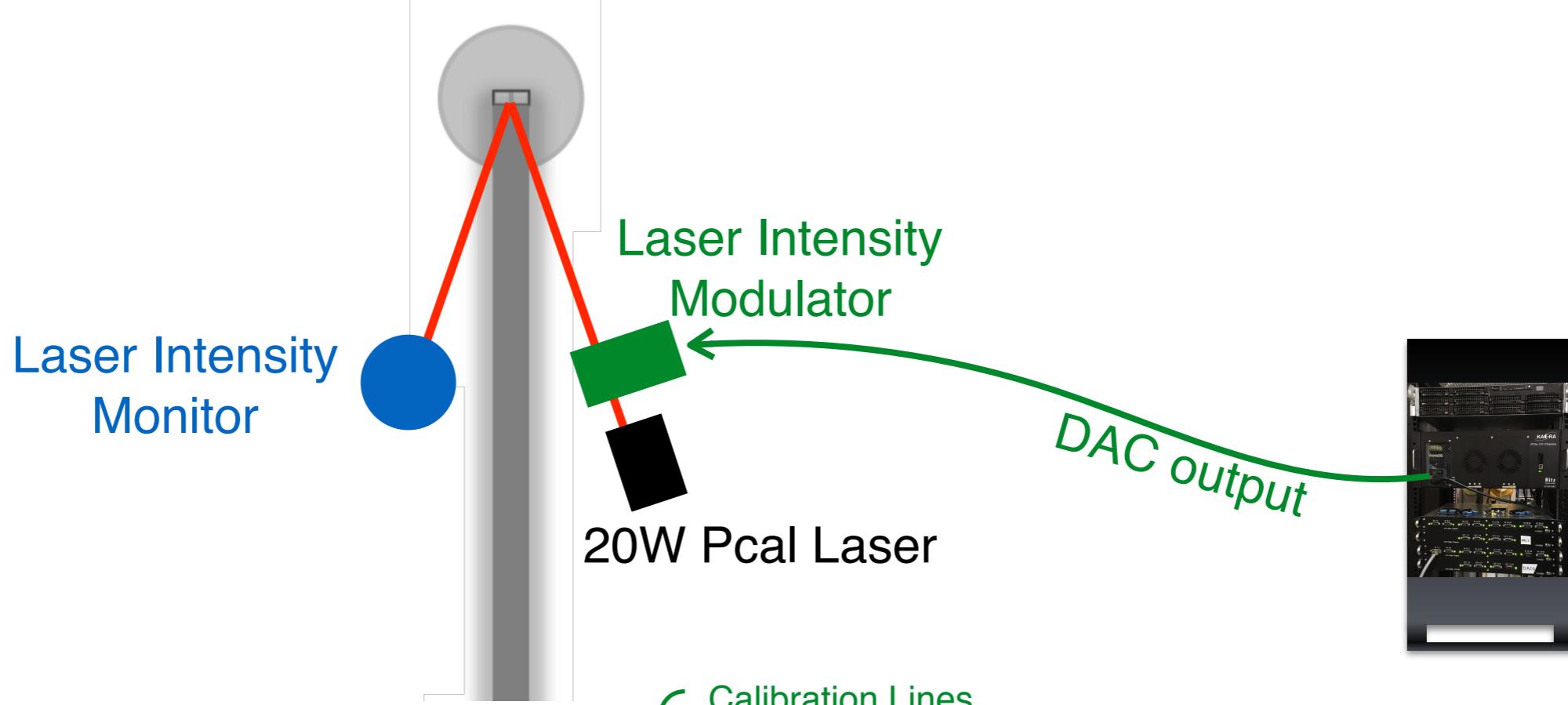
Radiation Force exerted by Pcal

$M(2\pi f)^2$

Force to length transfer function

$\Delta h(f)L$

strain sensitivity curve of KAGRA



Calibration Lines
 single frequency sine wave as standard length
 Theoretical GW waveform
 Binary Neutron Stars, Binary Blackholes
 Swept-Sine Excitation
 Transfer Function Measurement

~~Noise~~ from Control System, DAC, Power supply.....

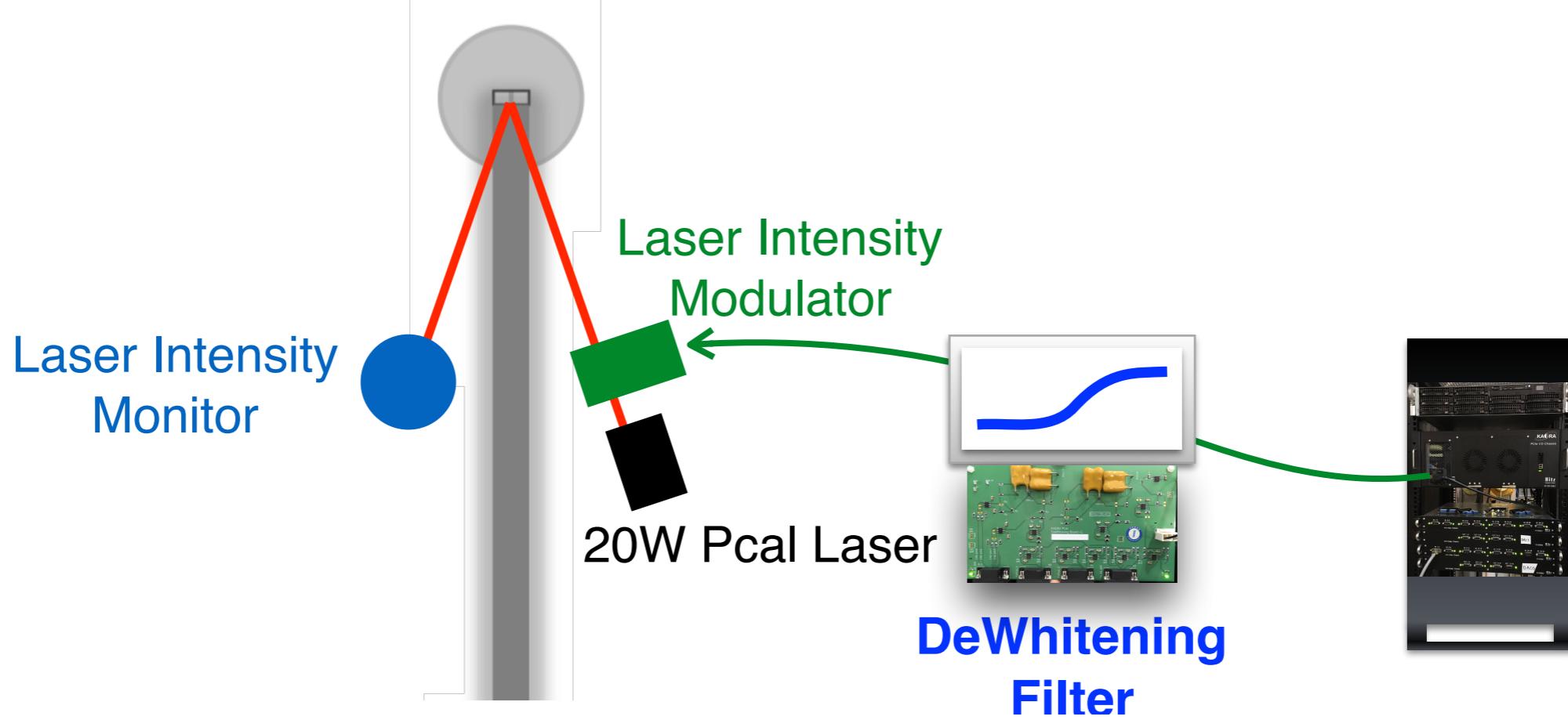
$$\Delta L(f) = \frac{2\Delta P(f) \cos(\theta)}{c} \frac{1}{M(2\pi f)^2} < \frac{1}{10} \Delta h(f)L$$

\nearrow
 Interferometer Mirror displacement
 induced by Pcal Laser intensity
 modulation

$\underbrace{\qquad\qquad\qquad}_{\text{Radiation Force exerted by Pcal}}$

$\underbrace{\qquad\qquad\qquad}_{\text{Force to length transfer function}}$

\nearrow
 strain sensitivity curve
 of KAGRA



Pole: 1kHz
Zero: 100Hz

$$\Delta L(f) = \underbrace{\frac{2\Delta P(f) \cos(\theta)}{c}}_{\text{Interferometer Mirror displacement induced by Pcal Laser intensity modulation}} \underbrace{\frac{1}{M(2\pi f)^2}}_{\text{Radiation Force exerted by Pcal}} < \frac{1}{10} \underbrace{\Delta h(f)L}_{\text{Force to length transfer function}}$$

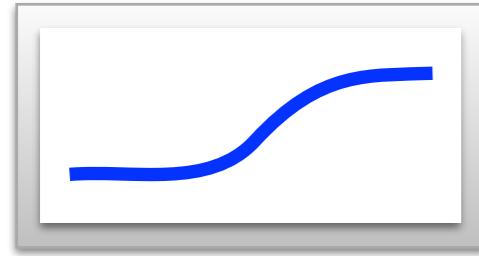
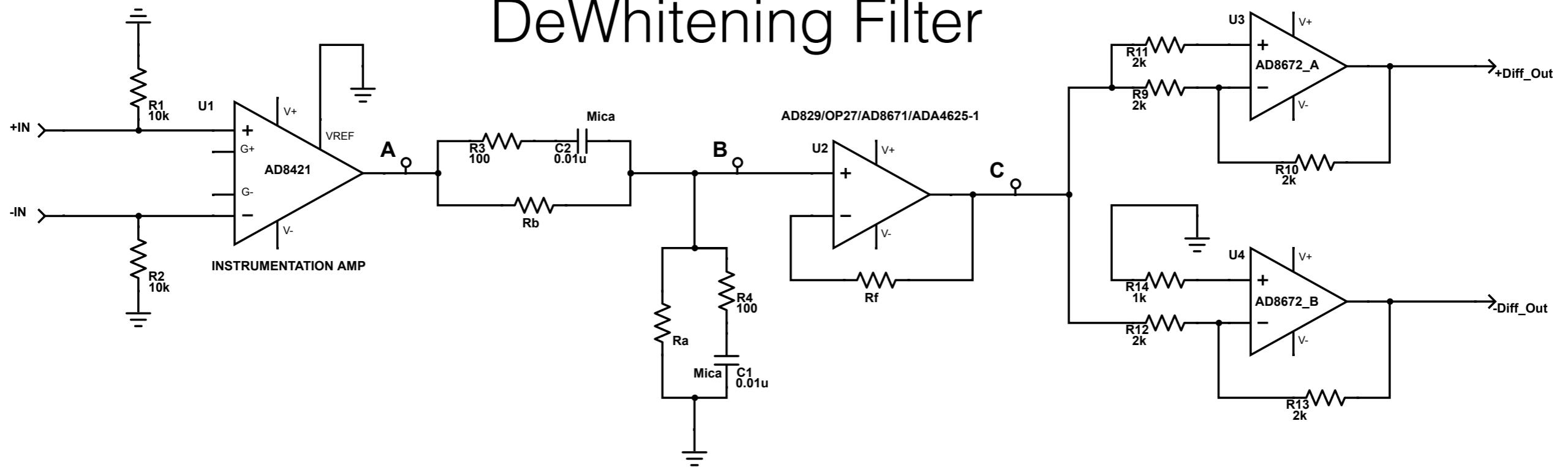
Interferometer Mirror displacement induced by Pcal Laser intensity modulation

Radiation Force exerted by Pcal

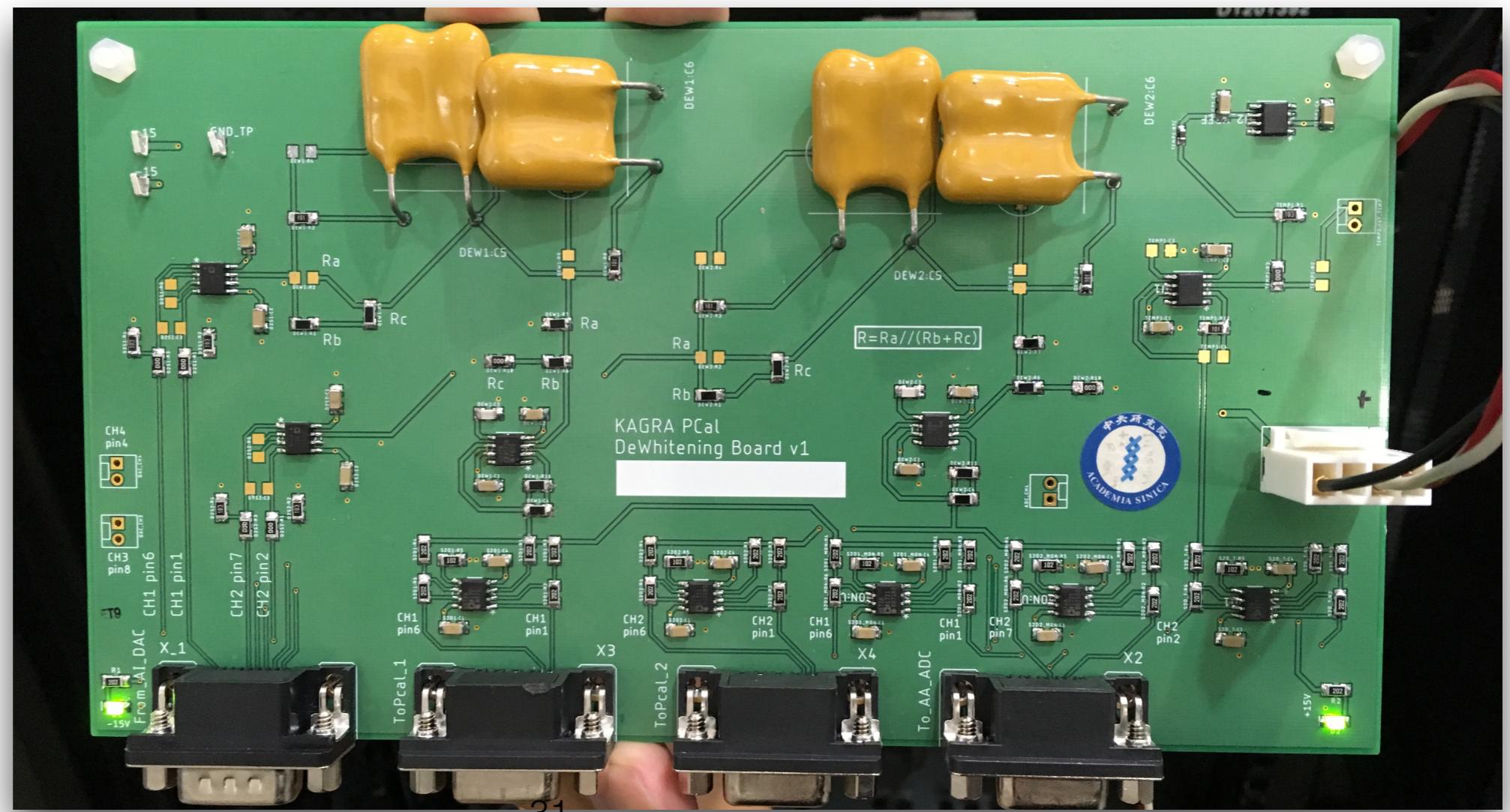
Force to length transfer function

strain sensitivity curve of KAGRA

DeWhitening Filter



Pole: 1kHz
Zero: 100Hz



Measurement of Noise Reduction Performance of De-Whitening Filter

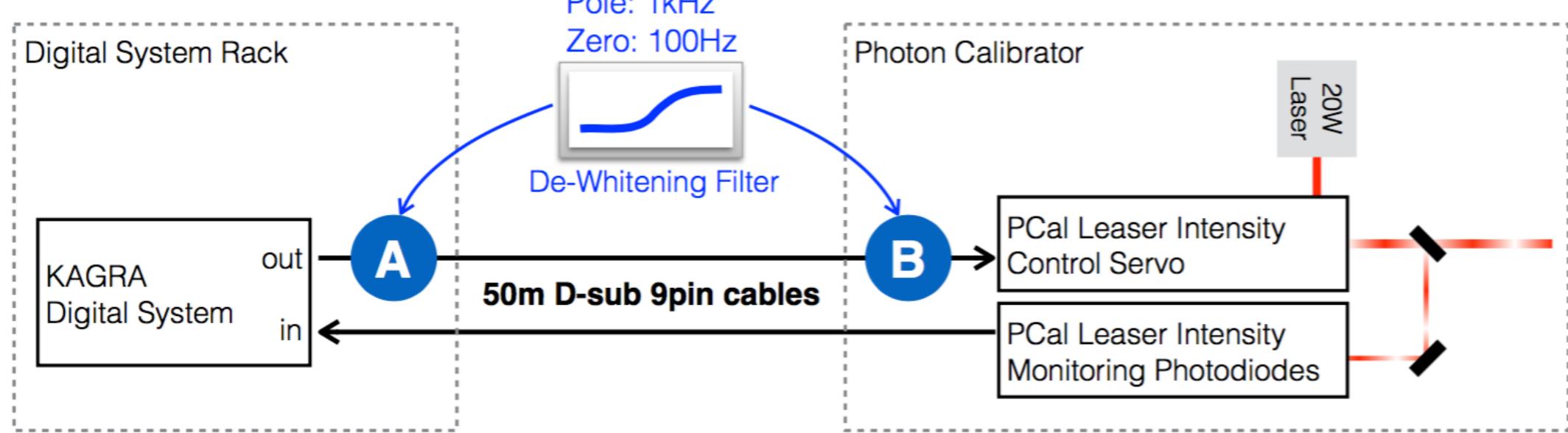


Figure 4.5: In order to reduce the noise coming from the digital system, the De-Whitening filter can be installed at either place A or B.

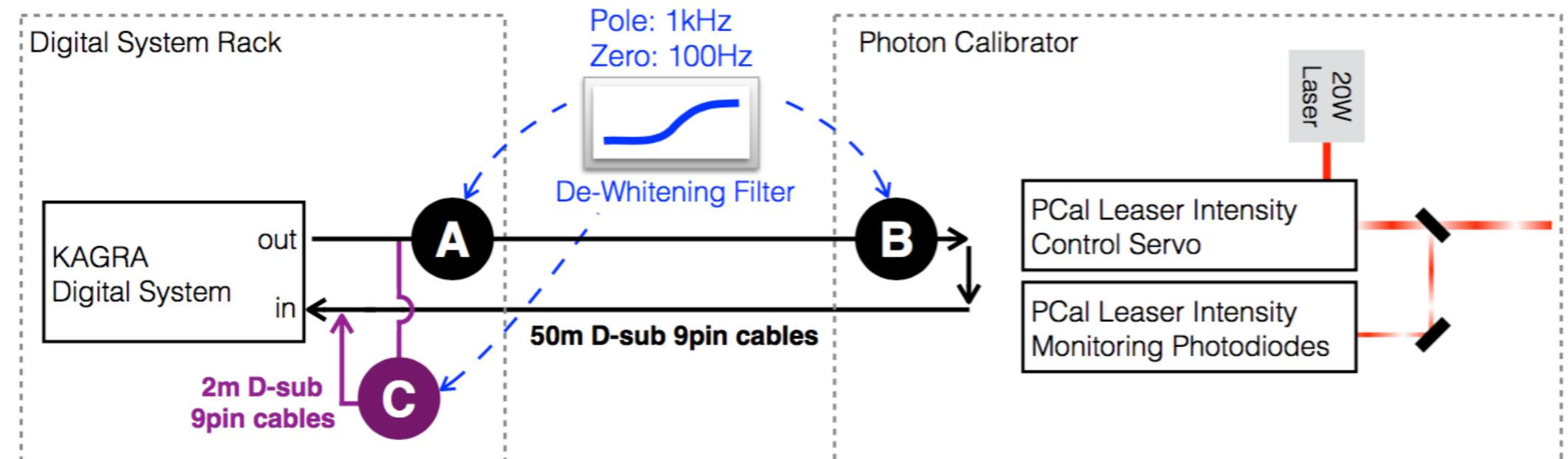


Figure 4.6: For our reference, we also measured the noise from the digital system without passing through the control loop of PCal. Place C means we connect De-Whitening filter in digital system rack with 2m cable only in order to investigate the influence from 50m cable.

KAGRA X-END 2m Cable

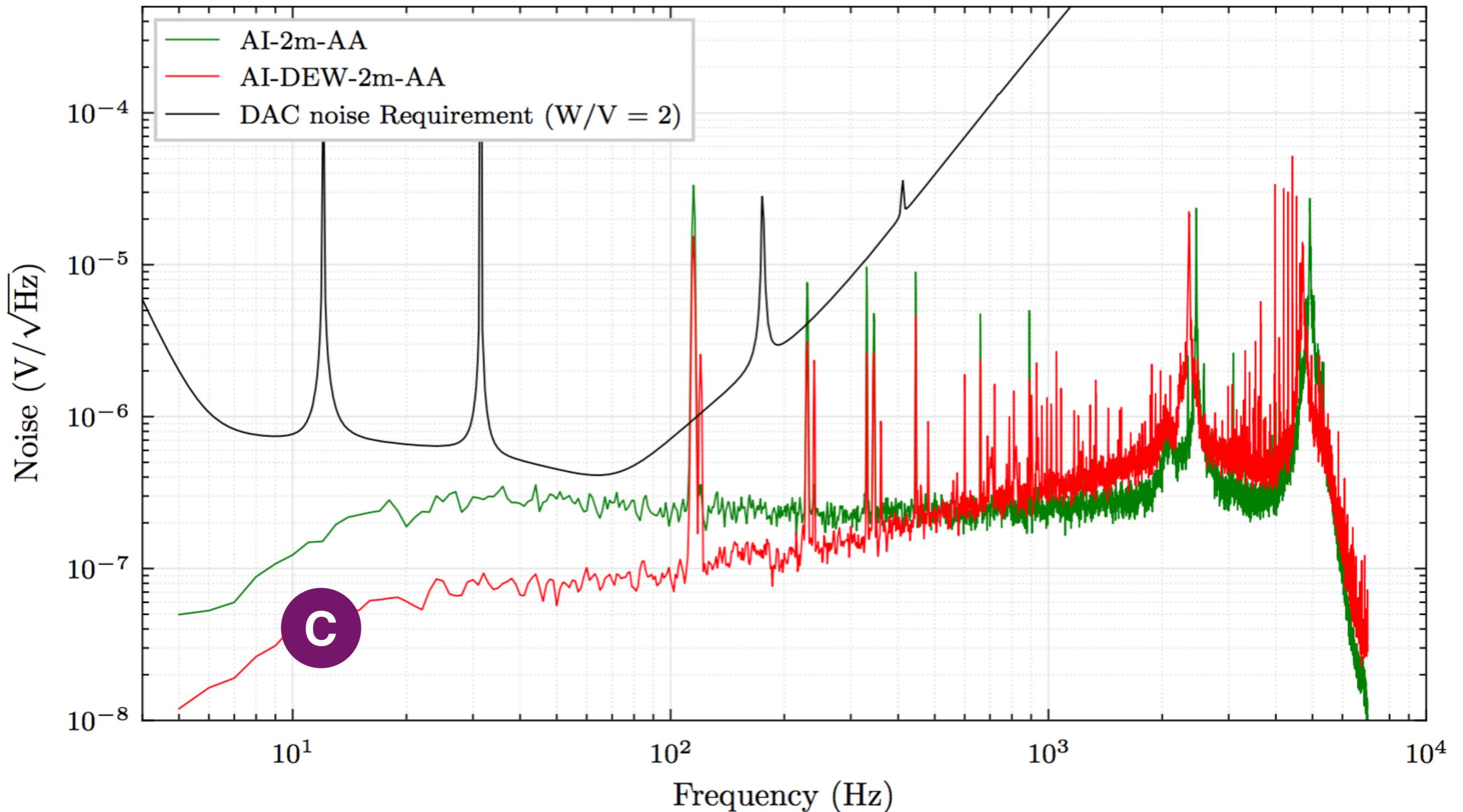


Figure 4.7: De-Whitening filter noise with short (2m) cable. The green line is the noise without De-Whitening filter, while the red one is the noise with De-Whitening filter.

KAGRA X-END 50m Cable

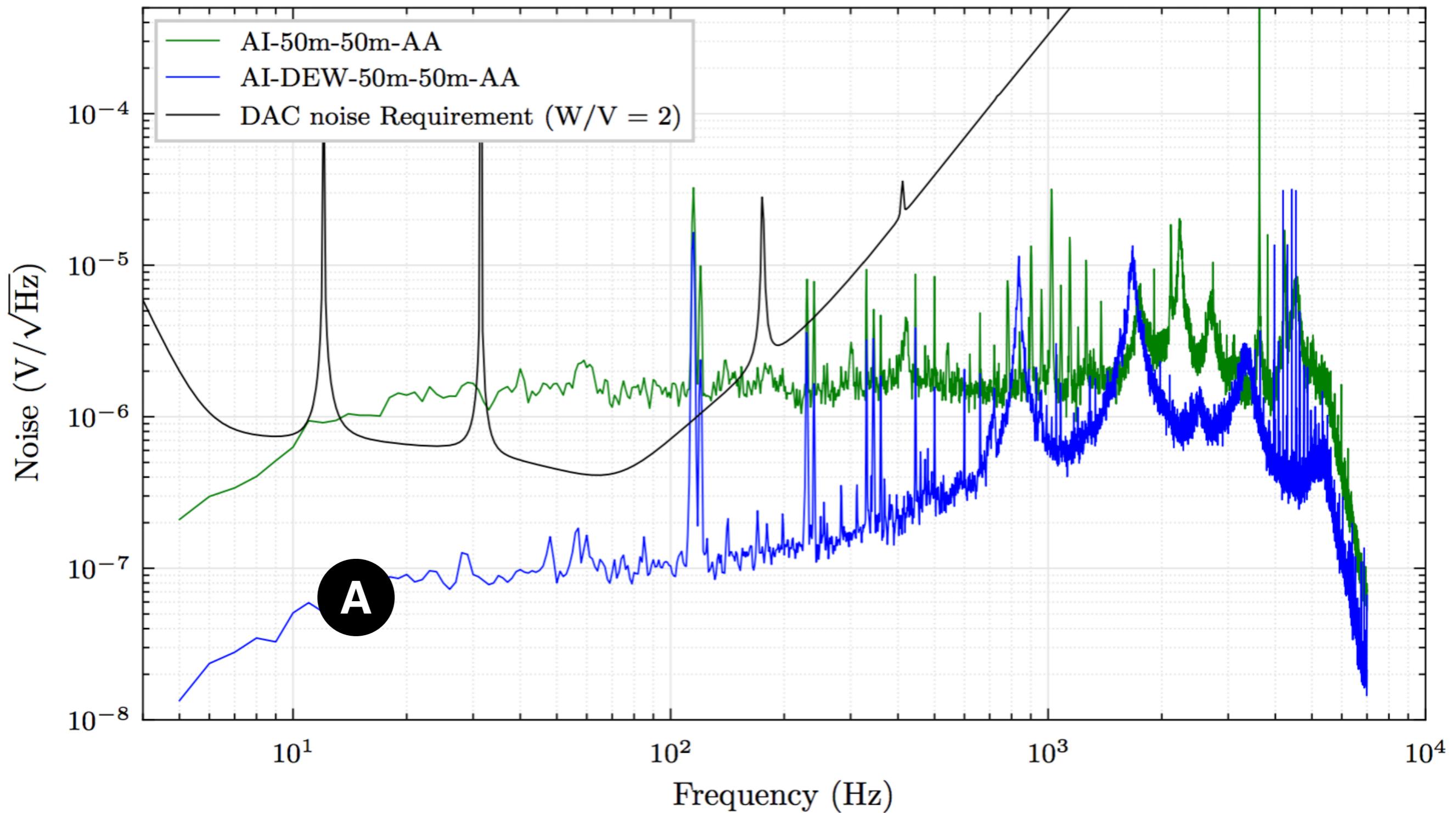


Figure 4.8: De-Whitening filter noise with long (50m) cable.

KAGRA X-END Different Cable Length

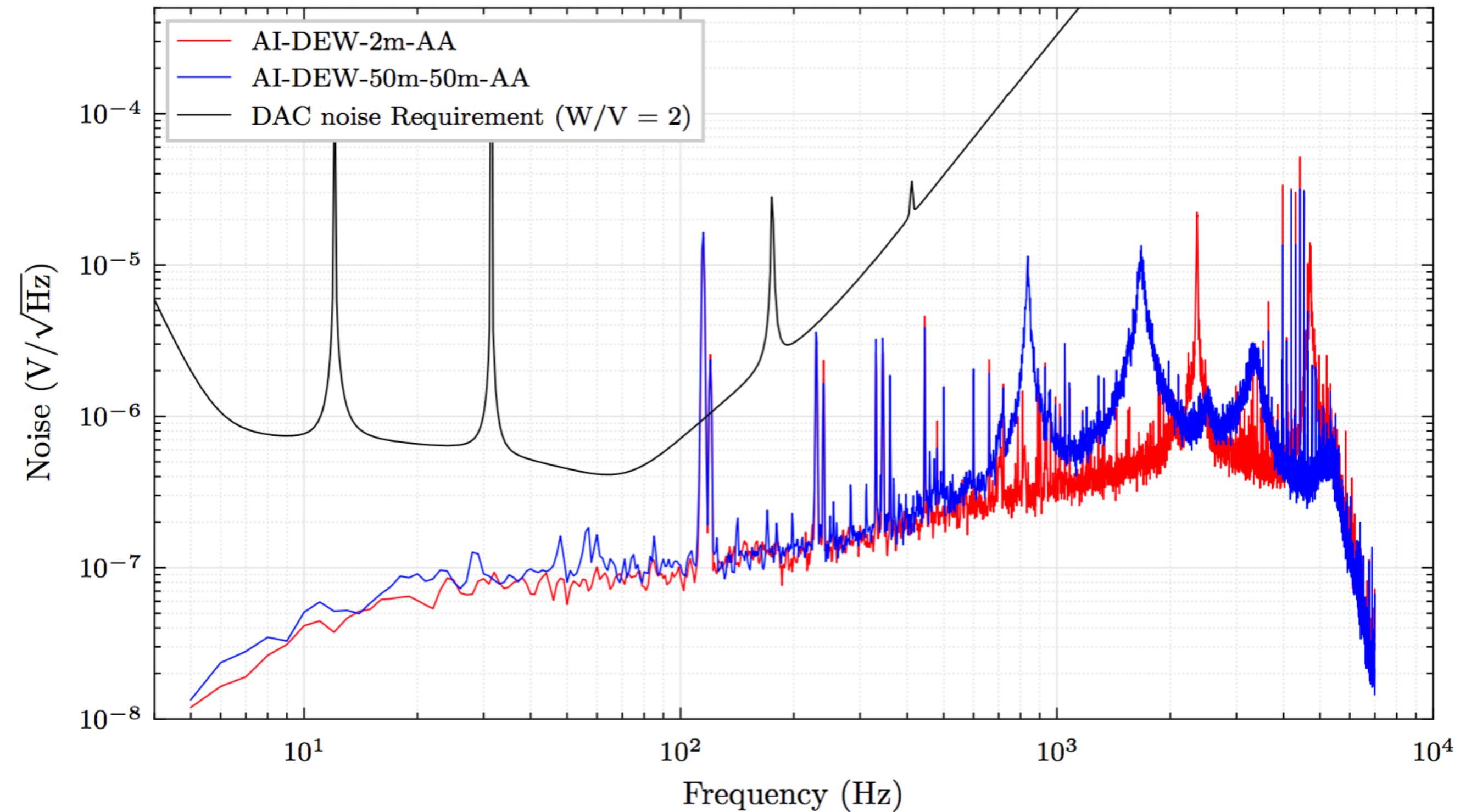


Figure 4.9: De-Whitening filter noise with different cable configuration.

KAGRA X-END DEW with Different Power Source

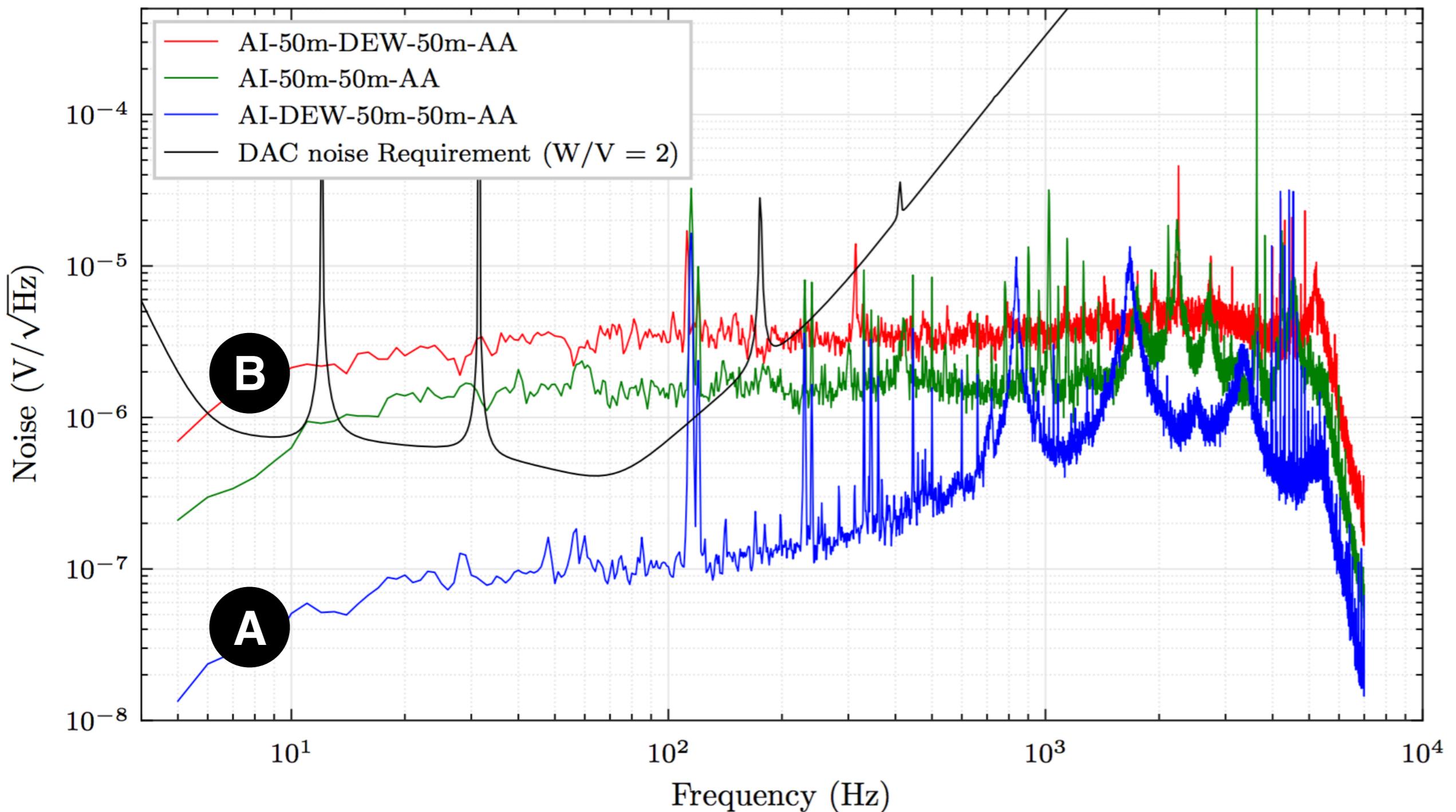


Figure 4.10: Noise measurement when the De-Whitening filter is installed at different location.

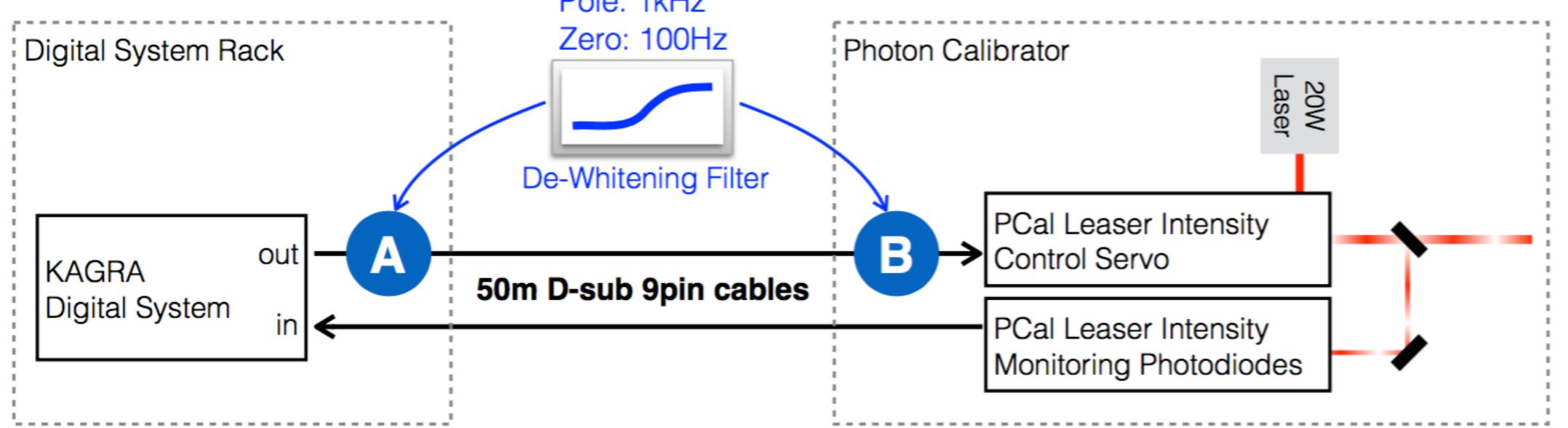


Figure 4.5: In order to reduce the noise coming from the digital system, the De-Whitening filter can be installed at either place A or B.

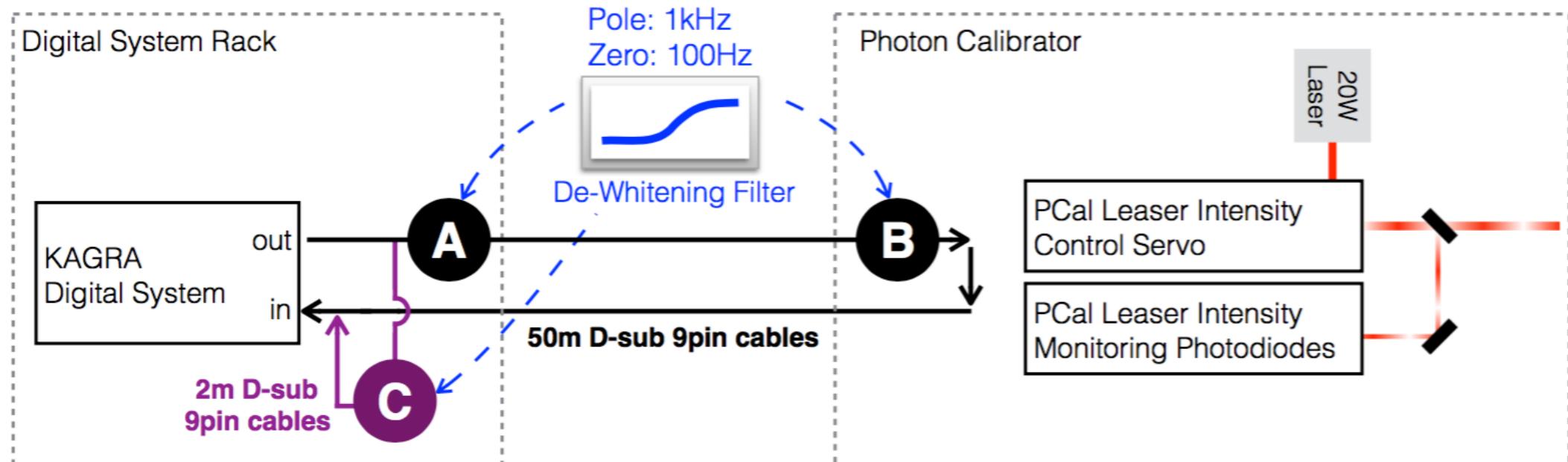


Figure 4.6: For our reference, we also measured the noise from the digital system without passing through the control loop of PCal. Place C means we connect De-Whitening filter in digital system rack with 2m cable only in order to investigate the influence from 50m cable.

KAGRA X-END

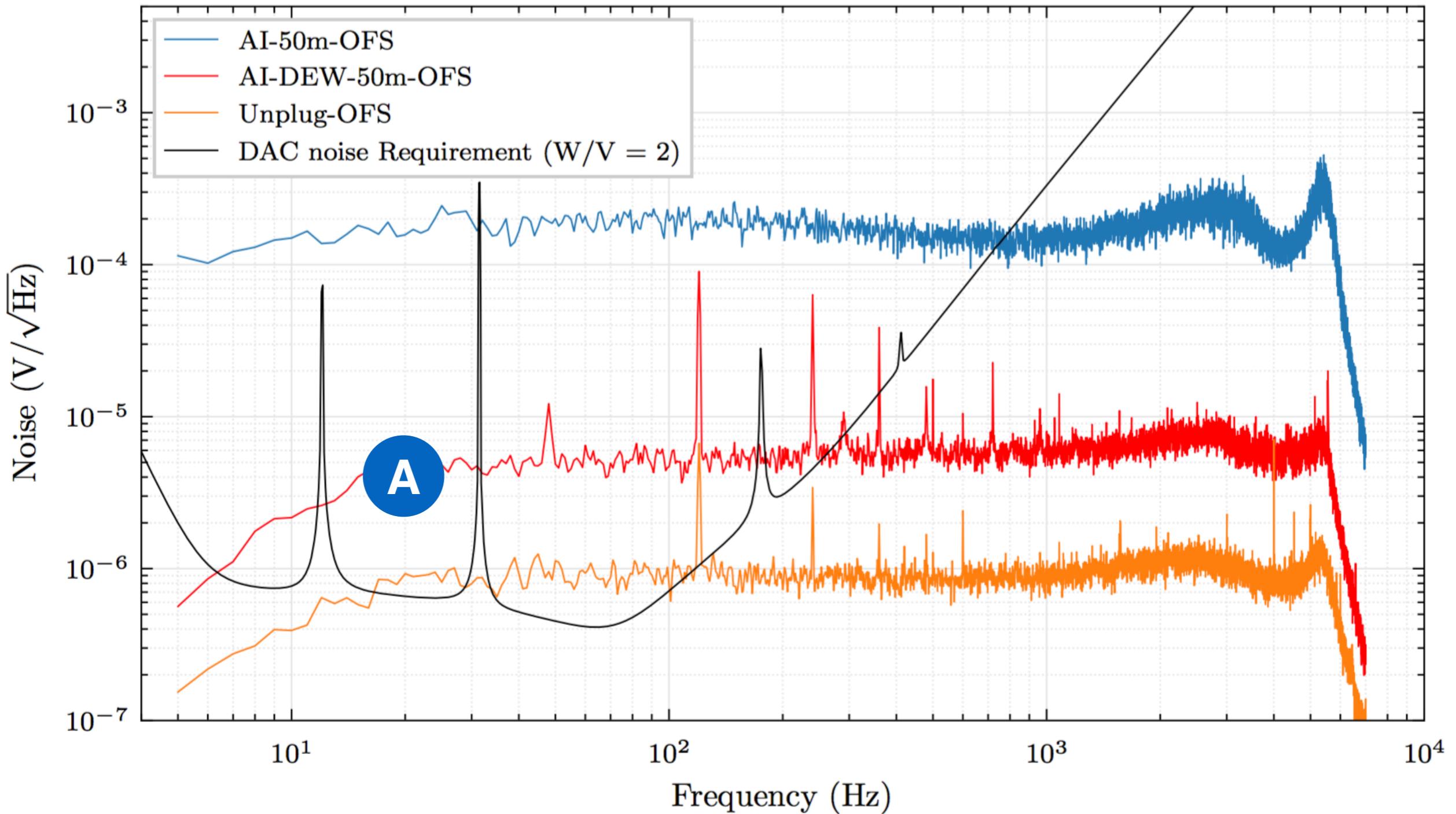


Figure 4.11: Noise measurement with PCal. These noises can be considered as laser intensity noises since we are measuring photodiode readout as depicted in Fig. 4.5. The blue line is the case without De-Whitening filer, while the red line is the case when De-Whitening filer has been installed at place A in Fig. 4.5. The orange line is measured when we disconnect our signal cable from the Laser Intensity Control Servo input port.

KAGRA X-END

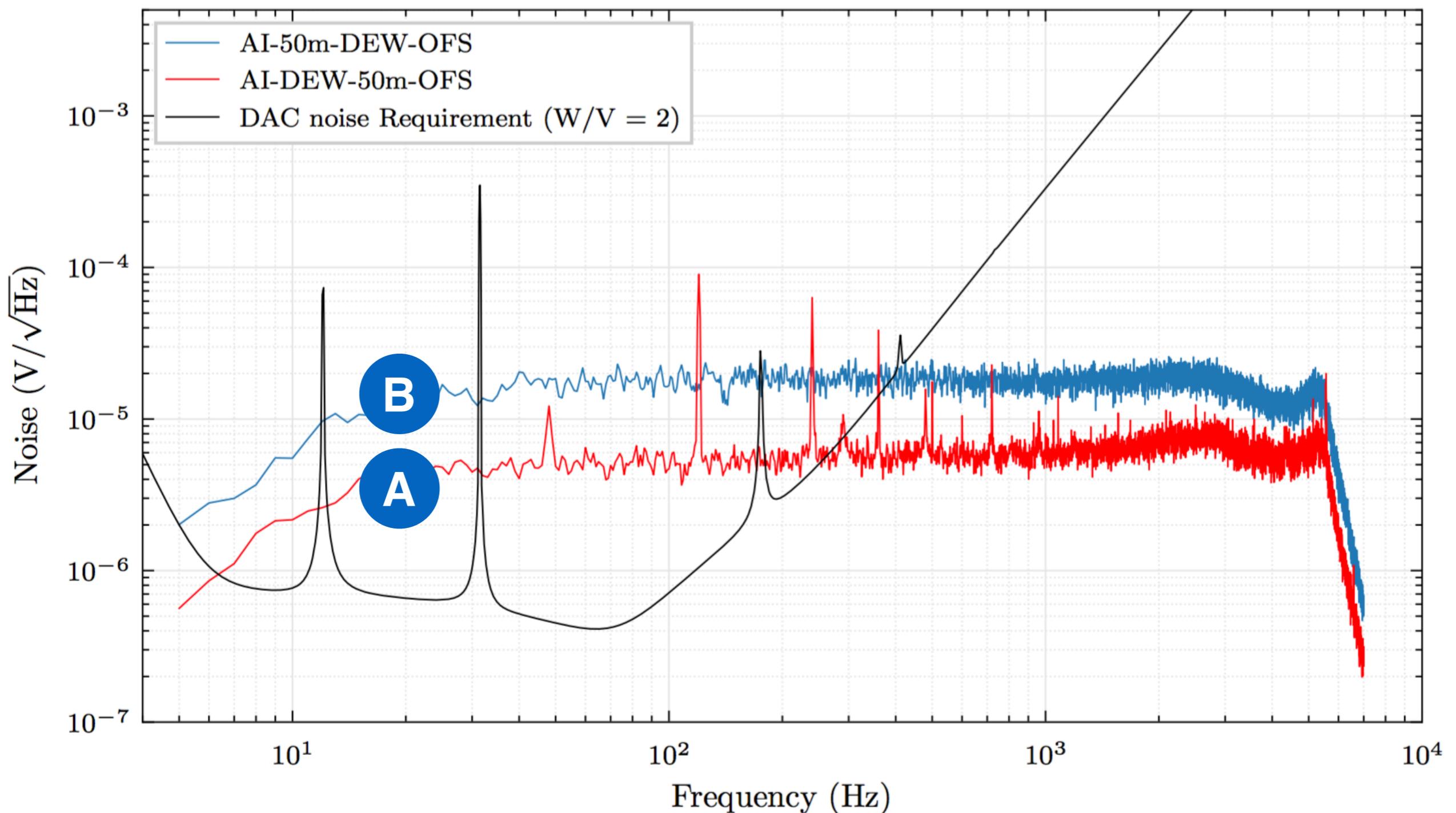


Figure 4.12: The red line and the blue line are measured when De-whitening is located at Place A and Place B in Fig. 4.5 respectively.

KAGRA X-END

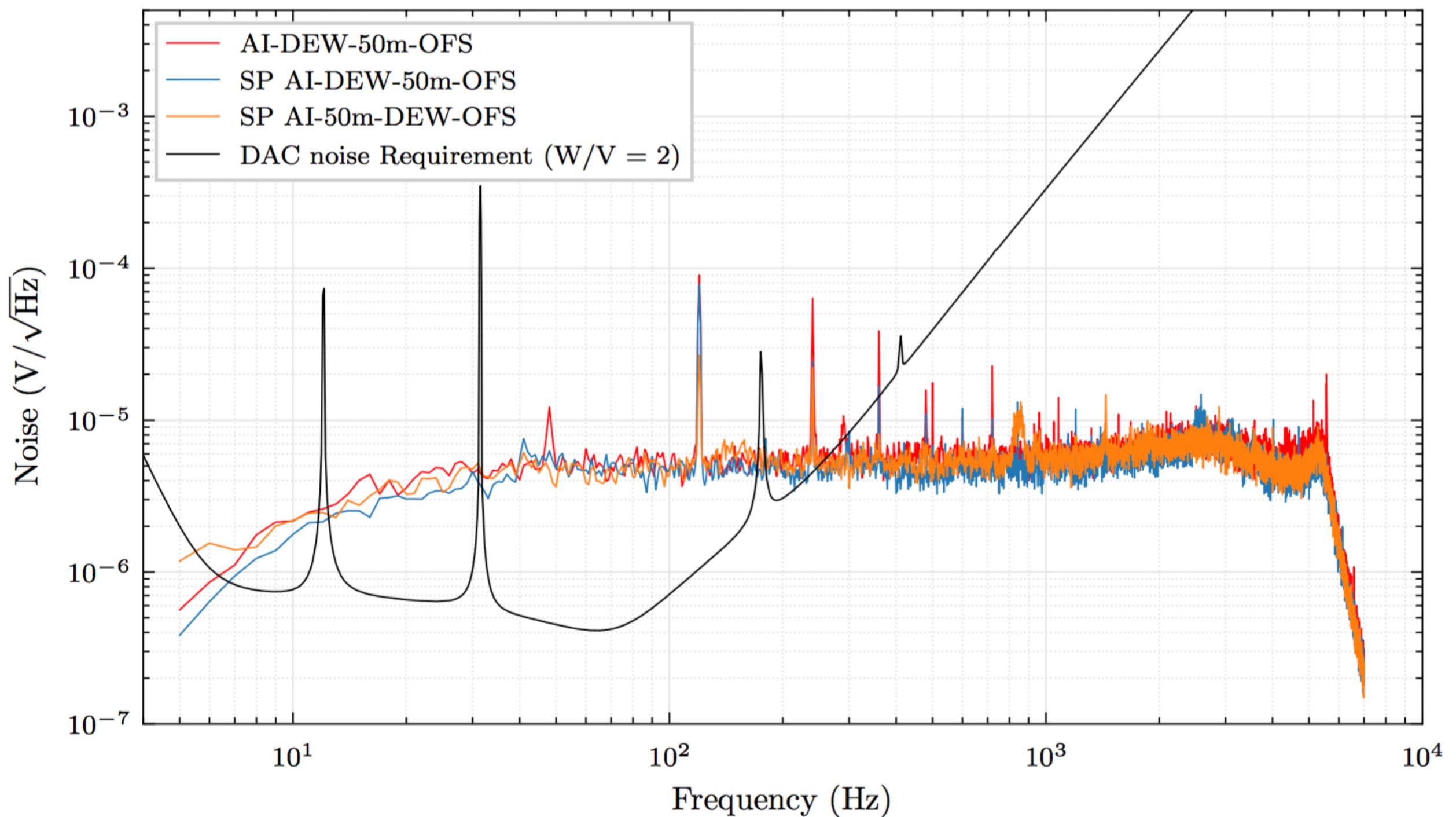
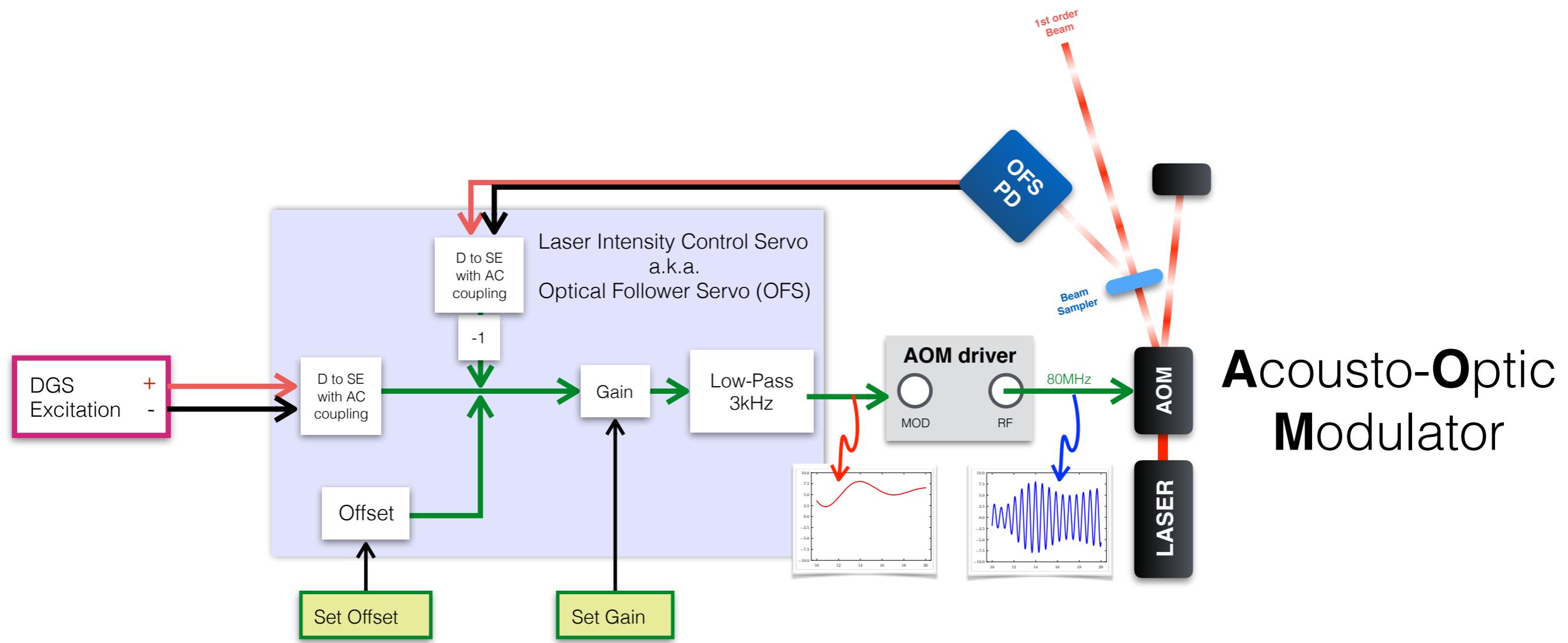
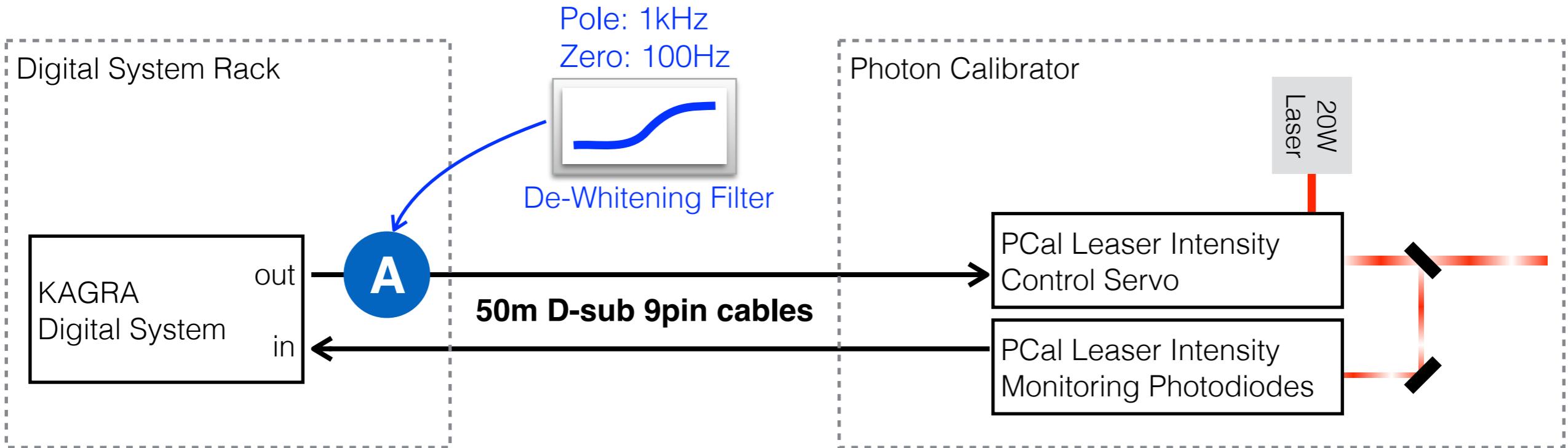
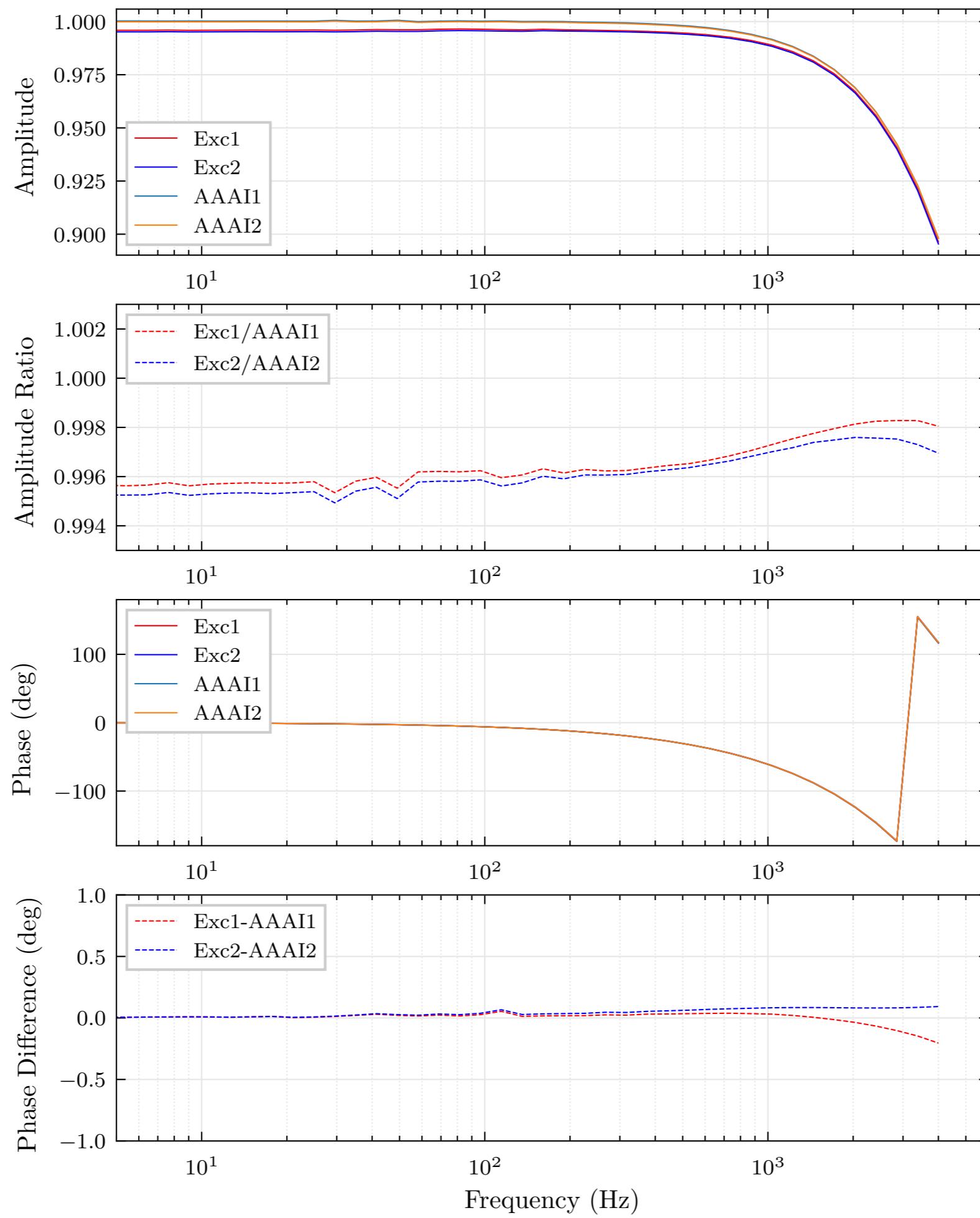


Figure 4.13: Lines labeled with “SP” were measured when we supplied digital system, De-Whitening filter and PCal with Same Power source located in digital system rack.

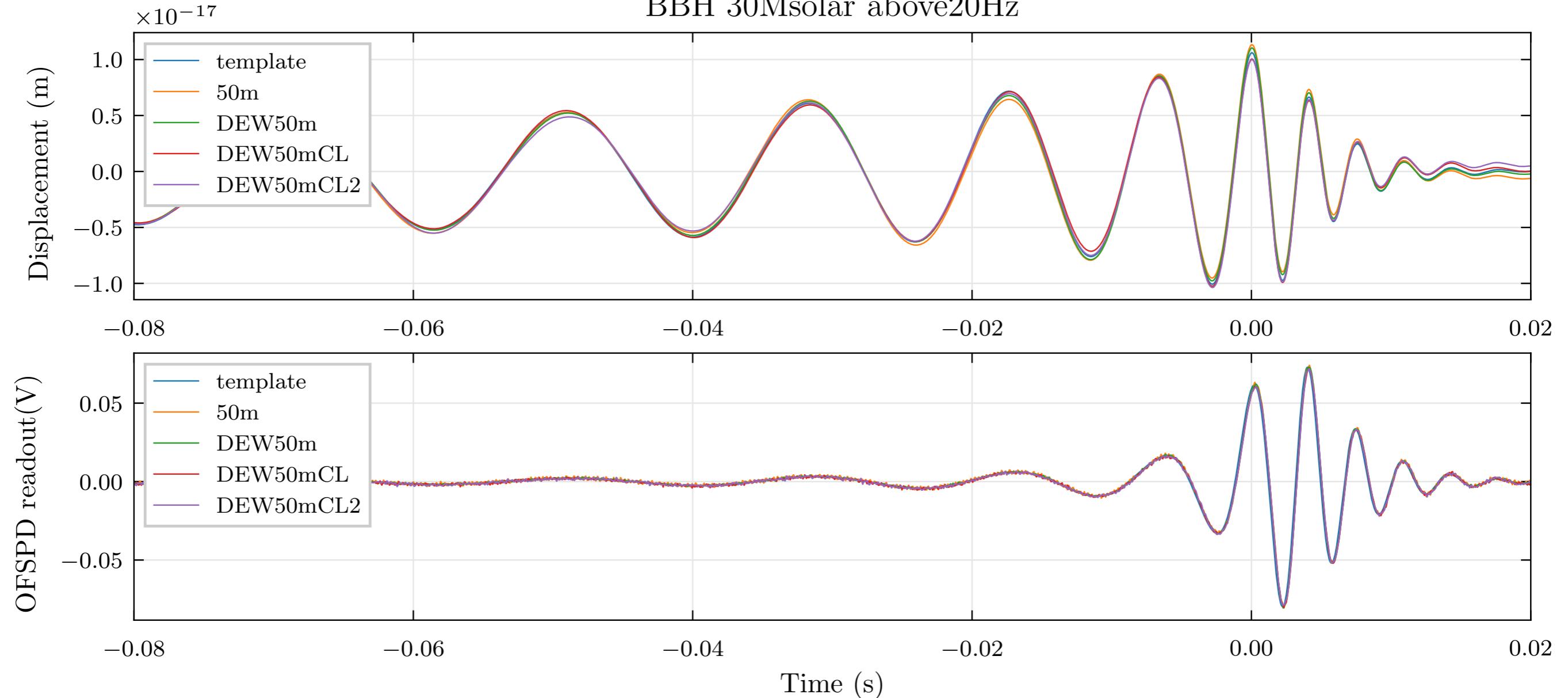
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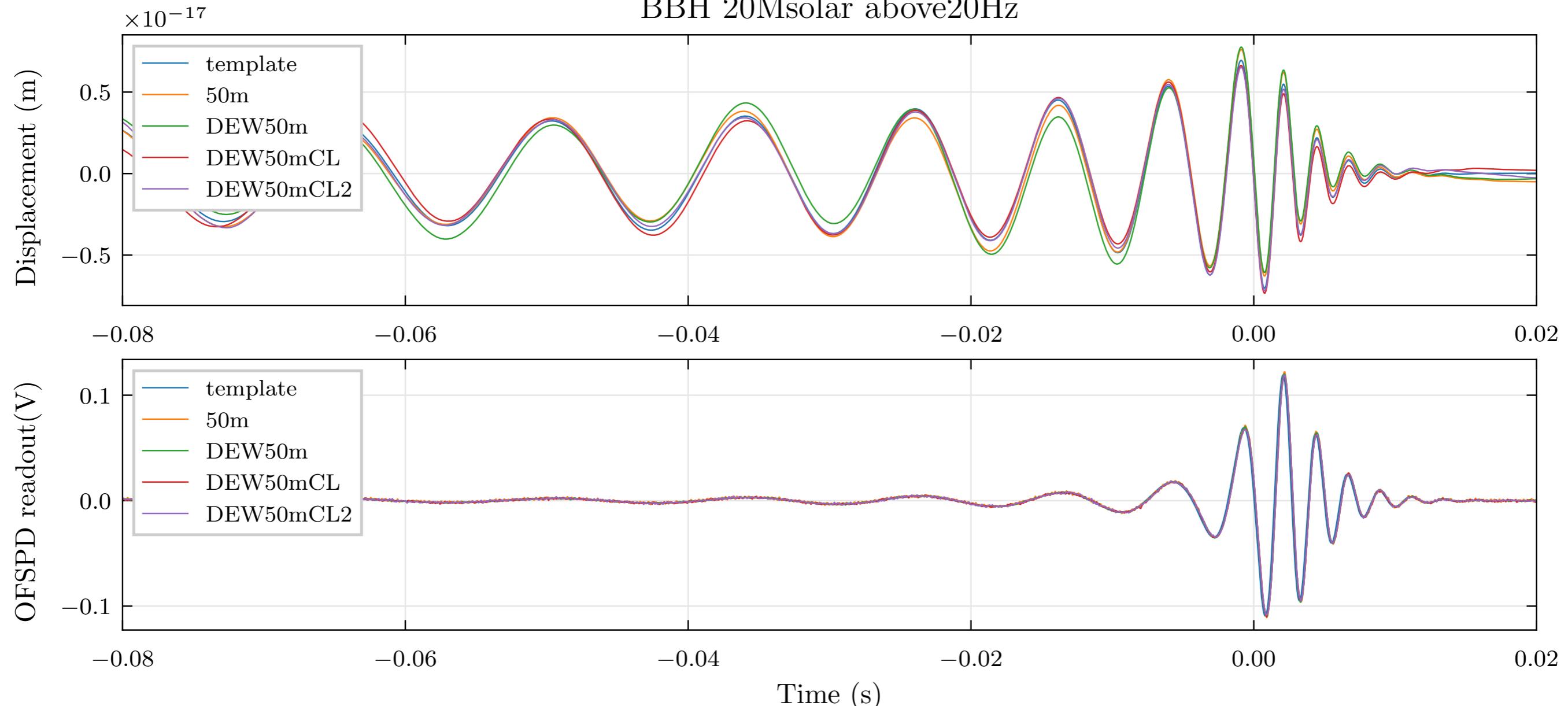




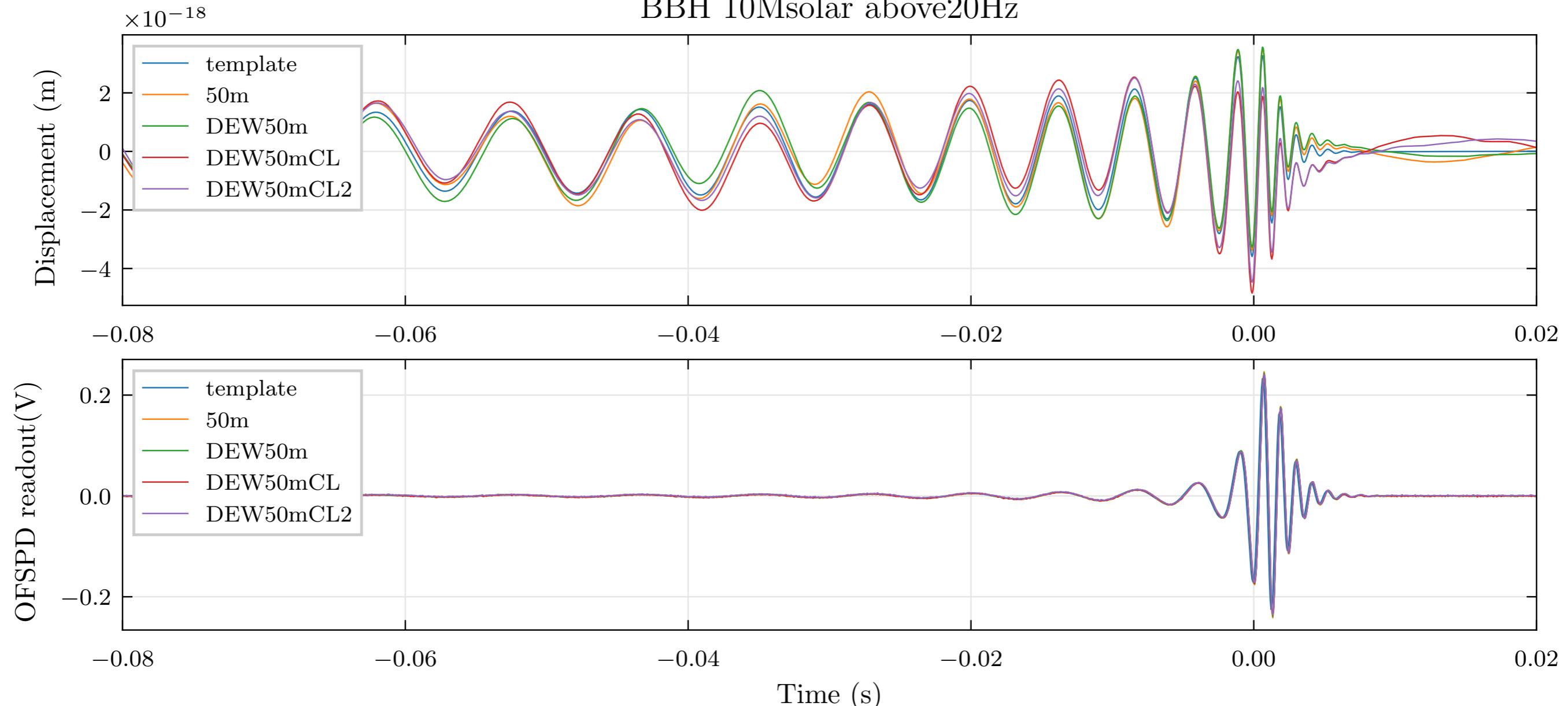
BBH 30Msolar above20Hz



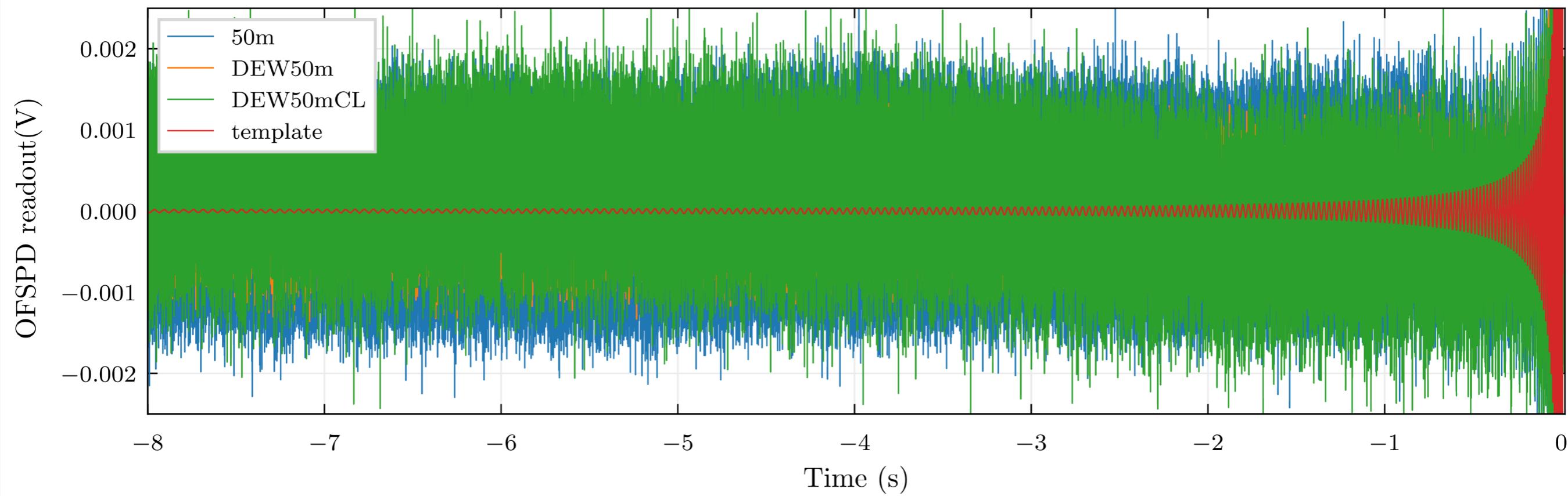
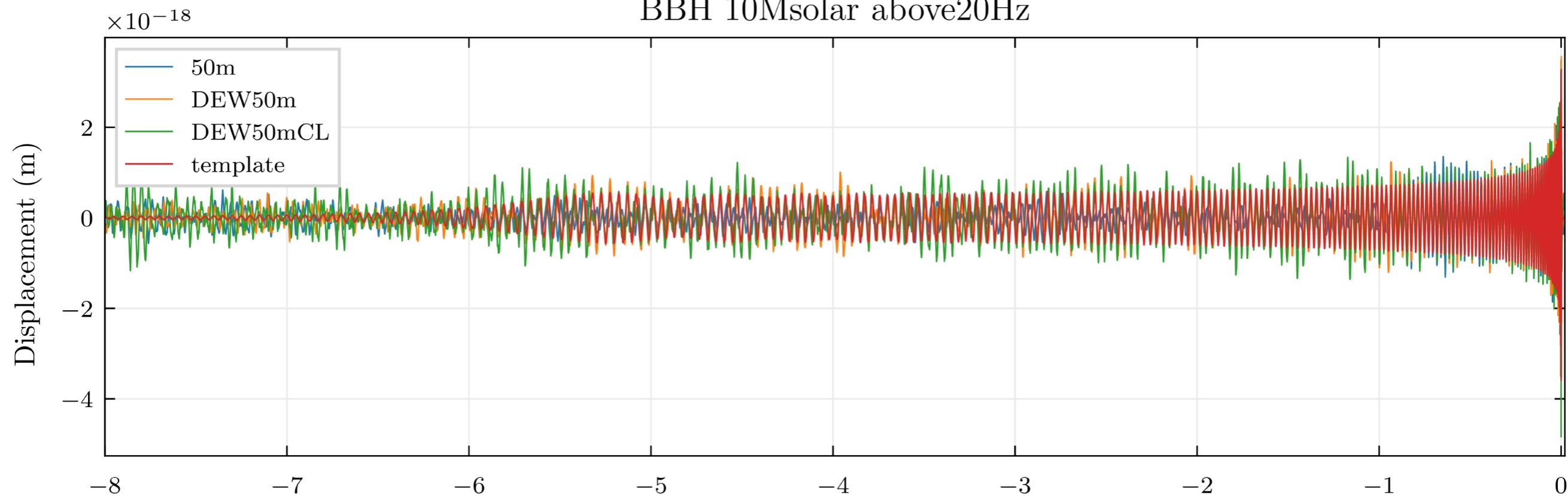
BBH 20Msolar above20Hz



BBH 10Msolar above20Hz



BBH 10Msolar above20Hz



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Future Works

- Performing Hardware Injection test using PCal with DeWhitening filter.
- Investigate analog excitation signal noise problem with long cable connection.

- Intrinsic limit of the De-Whitening Filter
- Noise shaping

END