

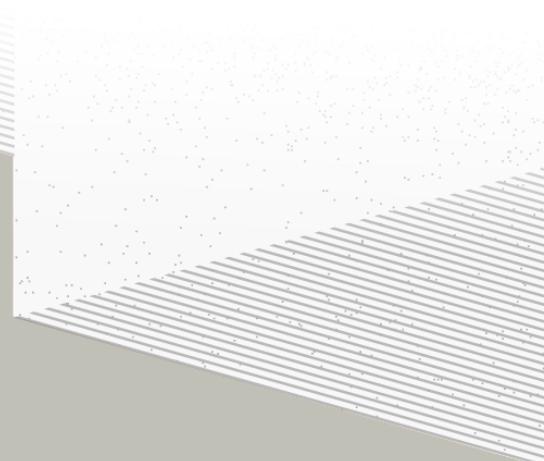
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# THÈSE DE DOCTORAT DE

L'UNIVERSITÉ DE NANTES

ÉCOLE DOCTORALE N°596  
*Matière, Molécules, Matériaux*  
Spécialité : *Physique des particules*



Par

**Léonard Imbert**

**Precision measurement of solar neutrino oscillation parameters  
with the JUNO small PMTs system and test of the unitarity of the  
PMNS matrix**

Thèse présentée et soutenue à Nantes, le Too soon and too early at the same time  
Unité de recherche : Laboratoire SUBATECH, UMR 6457

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# <sup>3</sup> List of Abbreviations

<b>ACU</b>	Automatic Calibration Unit
<b>BDT</b>	Boosted Decision Tree
<b>CD</b>	Central Detector
<b>CLS</b>	Cable Loop System
<b>CNN</b>	Convolutional NN
<b>DN</b>	Dark Noise
<b>DNN</b>	Deep NN
<b>FCDNN</b>	Fully Connected Deep NN
<b>GNN</b>	Graph NN
<b>GT</b>	Guiding Tube
<b>IBD</b>	Inverse Beta Decay
<b>IO</b>	Inverse Ordering
<b>JUNO</b>	Jiangmen Underground Neutrino Observatory
<b>LPMT</b>	Large PMT
<b>LS</b>	Liquid Scintillator
<b>MSE</b>	Mean Squared Error
<b>MC</b>	Monte Carlo simulation
<b>ML</b>	Machine Learning
<b>NMO</b>	Neutrino Mass Ordering
<b>NN</b>	Neural Network
<b>NO</b>	Normal Ordering
<b>NPE</b>	Number of Photo Electron
<b>OSIRIS</b>	Online Scintillator Internal Radioactivity Investigation System
<b>PE</b>	Photo Electron
<b>PMT</b>	Photo-Multipliers Tubes
<b>PRelu</b>	Parametrized Rectified Linear Unit
<b>ROV</b>	Remotely Operated under-LS Vehicle
<b>ReLU</b>	Rectified Linear Unit
<b>SPMT</b>	Small PMT
<b>TAO</b>	Taishan Antineutrino Observatory
<b>TR Area</b>	Total Reflexion Area
<b>TTS</b>	Time Transit Spread
<b>TT</b>	Top Tracker
<b>WCD</b>	Water Cherenkov Detector



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# Remerciements

<sup>59</sup>



<sup>60</sup> **Introduction**



<sup>61</sup> **Chapter 1**

<sup>62</sup> **Neutrino physics**

<sup>63</sup>

*The neutrino, or  $\nu$  for the close friends, a fascinating and invisible particle. Some will say that dark matter also have those property but at least we are pretty confident that neutrinos exists.*

<sup>64</sup> **1.1 Standard model**

<sup>65</sup> **1.1.1 Limits of the standard model**

<sup>66</sup> **1.2 Historic of the neutrino**

<sup>67</sup> **First theories**

<sup>68</sup> **Discovery**

<sup>69</sup> **Milestones and anomalies**

<sup>70</sup> **1.3 Oscillation**

<sup>71</sup> **1.3.1 Phenomologies**

<sup>72</sup> **1.4 Open questions**

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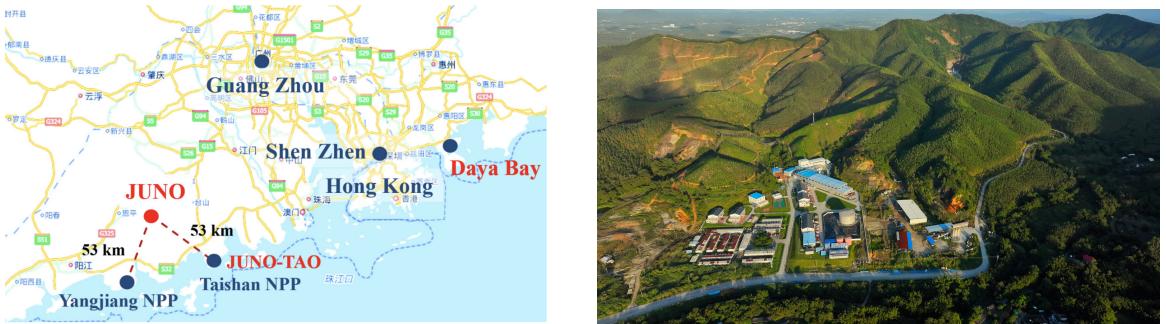
<sup>73</sup> **Chapter 2**

<sup>74</sup> **The JUNO experiment**

<sup>75</sup> *"Ave Juno, rosae rosam, et spiritus rex". It means nothing but I found it in tone.*

<sup>76</sup> The first idea of a medium baseline ( $\sim 52$  km) experiment, was explored in 2008 [1] where it was  
<sup>77</sup> demonstrated that the Neutrino Mass Ordering (NMO) could be determined by a medium baseline  
<sup>78</sup> experiment if  $\sin^2(2\theta_{13}) > 0.005$  without the requirements of accurate knowledge of the reactor  
<sup>79</sup> antineutrino spectra and the value of  $\Delta m_{32}^2$ . From this idea is born the Jiangmen Underground  
<sup>80</sup> Neutrino Observatory (JUNO) experiment.

<sup>81</sup> JUNO is a neutrino detection experiment under construction located in China, in Guangdong prov-  
<sup>82</sup> ing, near the city of Kaiping. Its main objectives are the determination of the mass ordering at the  
<sup>83</sup>  $3\text{-}4\sigma$  level in 6 years of data taking and the measurement at the sub-percent precision of the oscillation  
<sup>84</sup> parameters  $\Delta m_{21}^2$ ,  $\sin^2 \theta_{12}$ ,  $\Delta m_{32}^2$  and with less precision  $\sin^2 \theta_{13}$ [2].



<sup>85</sup> FIGURE 2.1 – **On the left:** Location of the JUNO experiment and its reactor sources in  
southern China. **On the right:** Aerial view of the experimental site

<sup>86</sup> For this JUNO will measure the electronic anti-neutrinos ( $\bar{\nu}_e$ ) flux coming from the nuclear reactors  
<sup>87</sup> of Taishan, Yangjiang, for a total power of  $26.6 \text{ GW}_{th}$ , and the Daya Bay power plant to a lesser  
<sup>88</sup> extent. All of those cores are the second-generation pressurized water reactors CPR1000, which is a  
<sup>89</sup> derivative of Framatome M310. Details about the power plants characteristics and their expected flux  
<sup>90</sup> of  $\bar{\nu}_e$  can be found in the table 2.1. The distance of 53 km has been specifically chosen to maximize  
the disappearance probability of the  $\bar{\nu}_e$ . The data taking is scheduled to start early 2025.

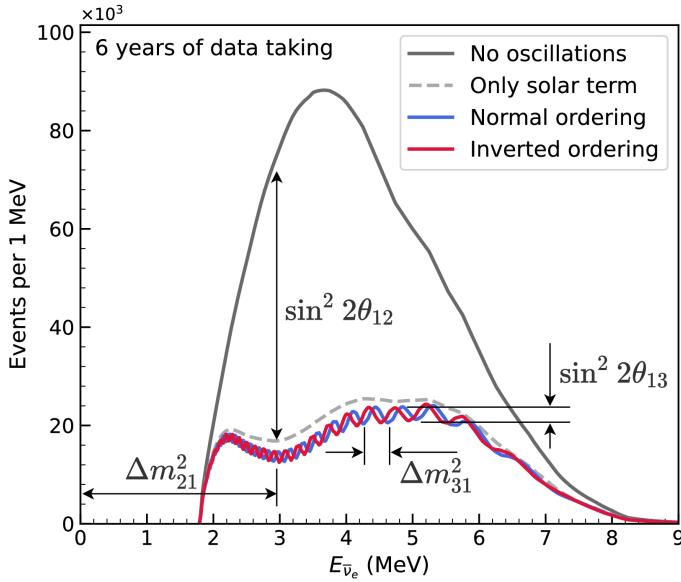


FIGURE 2.2 – Expected number of neutrinos event per MeV in JUNO after 6 years of data taking. The black curve shows the flux if there was no oscillation. The light gray curve shows the oscillation if only the solar terms are taken in account ( $\theta_{12}$ ,  $\Delta m_{21}^2$ ). The blue and red curve shows the spectrum in the case of, respectively, NO and IO. The dependency of the oscillation to the different parameters are schematized by the double sided arrows. We can see the NMO sensitivity by looking at the fine phase shift between the red and the blue curve.

## 91 2.1 Neutrinos physics in JUNO

92 Even if the JUNO design detailed in section 2.2 was optimized for the measurement of the NMO, its  
 93 large detection volume, excellent energy resolution and background level and understanding make it  
 94 also an excellent detector to measure the flux coming from other neutrino sources. Thus the scientific  
 95 program of JUNO extends way over reactor antineutrinos. The following section is an overview of  
 96 the different physics topic JUNO will contribute in the coming years.

### 97 2.1.1 Reactor neutrino oscillation for NMO and precise measurements

Previous works [1, 3] shows that oscillation parameters and the NMO can be observed by looking at the  $\bar{\nu}_e$  disappearance energy spectrum coming from medium baseline nuclear reactor. This disappearance probability can be expressed as [2] :

$$P(\bar{\nu}_e \rightarrow \bar{\nu}_e) = 1 - \sin^2 2\theta_{12} c_{13}^4 \sin^2 \frac{\Delta m_{21}^2 L}{4E} - \sin^2 2\theta_{13} \left[ c_{12}^2 \sin^2 \frac{\Delta m_{31}^2 L}{4E} + s_{12}^2 \sin^2 \frac{\Delta m_{32}^2 L}{4E} \right]$$

98 Where  $s_{ij} = \sin \theta_{ij}$ ,  $c_{ij} = \cos \theta_{ij}$ ,  $E$  is the  $\bar{\nu}_e$  energy and  $L$  is the baseline. We can see the sensitivity  
 99 to the NMO in the dependency to  $\Delta m_{32}^2$  and  $\Delta m_{31}^2$  causing a phase shift of the spectrum as we can  
 100 see in the figure 2.2. By carefully adjusting a theoretical spectrum to the data, one can extract the  
 101 NMO and the oscillation parameters. The statistic procedure used to adjust the theoretical spectrum  
 102 is reviewed in more details in the section 2.7. To reach the desired sensitivity, JUNO must meet  
 103 multiple requirements but most notably:

- 104     1. An energy resolution of  $3\%/\sqrt{E(\text{MeV})}$  to be able to distinguish the fine structure of the fast  
 105     oscillation.
- 106     2. An energy precision of 1% in order to not err on the location of the oscillation pattern.
- 107     3. A baseline of  $53 \pm 0.5$  km to maximise the  $\bar{\nu}_e$  oscillation probability.
- 108     4. At least  $\approx 100,000$  events to limit the spectrum distortion due to statistical uncertainties.

109  **$\bar{\nu}_e$  flux coming from nuclear power plants**

110 To get such high measurements precision, it is necessary to have a very good understanding of the  
 111 sources characteristics. For its NMO and precise measurement studies, JUNO will observe the energy  
 112 spectrum of neutrinos coming from the nuclear power plants Taishan and Yangjiang's cores, located  
 113 at 53 km of the detector to maximise the disappearance probability of the  $\bar{\nu}_e$ .

Reactor	Power (GW <sub>th</sub> )	Baseline (km)	IBD Rate (day <sup>-1</sup> )	Relative Flux (%)
Taishan	9.2	52.71	15.1	32.1
Core 1	4.6	52.77	7.5	16.0
Core 2	4.6	52.64	7.6	16.1
Yangjiang	17.4	52.46	29.0	61.5
Core 1	2.9	52.74	4.8	10.1
Core 2	2.9	52.82	4.7	10.1
Core 3	2.9	52.41	4.8	10.3
Core 4	2.9	52.49	4.8	10.2
Core 5	2.9	52.11	4.9	10.4
Core 6	2.9	52.19	4.9	10.4
Daya Bay	17.4	215	3.0	6.4

TABLE 2.1 – Characteristics of the nuclear power plants observed by JUNO. The IBD rate are estimated from the baselines, the reactors full thermal power, selection efficiency and the current knowledge of the oscillation parameters

114 The  $\bar{\nu}_e$  coming from reactors are emitted from  $\beta$ -decay of unstable fission fragments. The Taishan  
 115 and Yangjiang reactors are Pressurised Water Reactor (PWR), the same type as Daya Bay. In those  
 116 type of reactor more the 99.7 % and  $\bar{\nu}_e$  are produced by the fissions of four fuel isotopes  $^{235}\text{U}$ ,  $^{238}\text{U}$ ,  
 117  $^{239}\text{Pu}$  and  $^{241}\text{Pu}$ . The neutrino flux per fission of each isotope is determined by the inversion of the  
 118 measured  $\beta$  spectra of fission product [4–8] or by calculation using the nuclear databases [9, 10].

119 The neutrino flux coming from a reactor at a time  $t$  can be predicted using

$$\phi(E_\nu, t)_r = \frac{W_{th}(t)}{\sum_i f_i(t) e_i} \sum_i f_i(t) S_i(E_\nu) \quad (2.1)$$

120 where  $W_{th}(t)$  is the thermal power of the reactor,  $f_i(t)$  is the fraction fission of the  $i$ th isotope,  $e_i$  its  
 121 thermal energy released in each fission and  $S_i(e_\nu)$  the neutrino flux per fission for this isotope. Using  
 122 this method, the flux uncertainty is expected to be of an order of 2-3 % [11].

123 In addition to those prediction, a satellite experiment named TAO[12] will be setup near the reactor  
 124 core Taishan-1 to measure with an energy resolution of 2% at 1 MeV the neutrino flux coming from  
 125 the core, more details can be found in section 2.4.1. It will help identifying unknown fine structure  
 126 and give more insight on the  $\bar{\nu}_e$  flux coming from this reactor.

127 One the open issue about reactor anti-neutrinos flux is the so-called neutrino anomaly [13], an  
 128 unexpected surplus of neutrino emission in the spectra around 5 MeV. Multiples scientists are trying  
 129 to explain this surplus by advanced recalculation of the nuclei model during beta decay [14, 15] but  
 130 no consensus on this issue has been reached yet.

131 **Background in the neutrinos reactor spectrum**

132 Considering the close reactor neutrinos flux as the main signal, the signals that are considered as  
 133 background are:

- 134 — The geoneutrinos producing background in the  $0.511 \sim 2.7$  MeV region.
- 135 — The neutrinos coming from the other nuclear reactors around Earth.

136 In addition to all those physics signal, non-neutrinos signal that would mimic an IBD will also be  
 137 present. It is composed of:

- 138 — The signal coming from radioactive decay ( $\alpha$ ,  $\gamma$ ,  $\beta$ ) from natural radioactive isotopes in the  
 139 material of the detector.
- 140 — Cosmogenic event such as fast neutrons and activated isotopes induced by muons passing  
 141 through the detector, most notably the spallation on  $^{12}\text{C}$ .

142 All those events represent a non-negligable part of the spectrum as shown in figure 2.3.

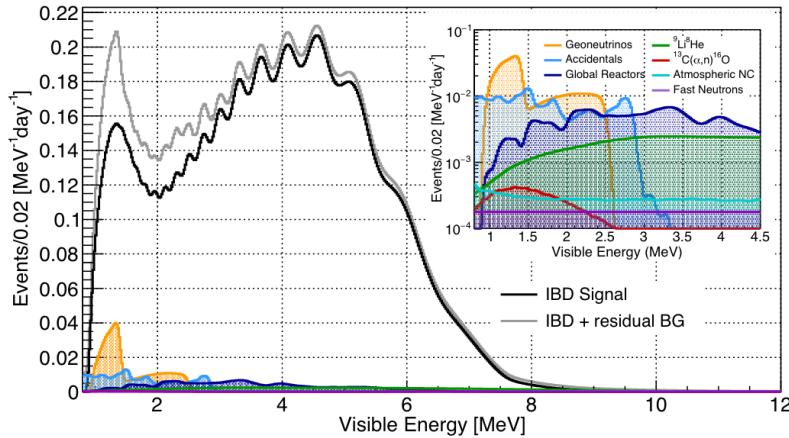


FIGURE 2.3 – Expected visible energy spectrum measured with the LPMT system with (grey) and without (black) backgrounds. The background amount for about 7% of the IBD candidate and are mostly localized below 3 MeV [11]

143 **Identification of the mass ordering**

144 To identify the mass ordering, we adjust the theoretical neutrino energy spectrum under the two  
 145 hypothesis of NO and IO. Those give us two  $\chi^2$ , respectively  $\chi^2_{\text{NO}}$  and  $\chi^2_{\text{IO}}$ . By computing the  
 146 difference  $\Delta\chi^2 = \chi^2_{\text{NO}} - \chi^2_{\text{IO}}$  we can determine the most probable mass ordering and the confidence  
 147 interval: NO if  $\Delta\chi^2 > 0$  and IO if  $\Delta\chi^2 < 0$ . Current studies shows that the expected sensitivity  
 148 the mass ordering would be of  $3.4\sigma$  after 6 years of data taking in nominal setup[2]. More detailed  
 149 explanations about the procedure can be found in the section 2.7.

150 **Precise measurement of the oscillations parameters**

151 The oscillations parameters  $\theta_{12}$ ,  $\theta_{13}$ ,  $\Delta m_{21}^2$ ,  $\Delta m_{31}^2$  are free parameters in the fit of the oscillation  
 152 spectrum. The precision on those parameters have been estimated and are shown in table 2.2. Wee  
 153 see that for  $\theta_{12}$ ,  $\Delta m_{21}^2$ ,  $\Delta m_{31}^2$ , precision at 6 years is better than the reference precision by an order of  
 154 magnitude [11]

	Central Value	PDG 2020	100 days	6 years	20 years
$\Delta m_{31}^2 (\times 10^{-3} \text{ eV}^2)$	2.5283	$\pm 0.034$ (1.3%)	$\pm 0.021$ (0.8%)	$\pm 0.0047$ (0.2%)	$\pm 0.0029$ (0.1%)
$\Delta m_{21}^2 (\times 10^{-3} \text{ eV}^2)$	7.53	$\pm 0.18$ (2.4%)	$\pm 0.074$ (1.0%)	$\pm 0.024$ (0.3%)	$\pm 0.017$ (0.2%)
$\sin^2 \theta_{12}$	0.307	$\pm 0.013$ (4.2%)	$\pm 0.0058$ (1.9%)	$\pm 0.0016$ (0.5%)	$\pm 0.0010$ (0.3%)
$\sin^2 \theta_{13}$	0.0218	$\pm 0.0007$ (3.2%)	$\pm 0.010$ (47.9%)	$\pm 0.0026$ (12.1%)	$\pm 0.0016$ (7.3%)

TABLE 2.2 – A summary of precision levels for the oscillation parameters. The reference value (PDG 2020 [16]) is compared with 100 days, 6 years and 20 years of JUNO data taking.

### 2.1.2 Other physics

While the design of JUNO is tailored to measure  $\bar{\nu}_e$  coming from nuclear reactor, JUNO will be able to detect neutrinos coming from other sources thus allowing for a wide range of physics studies as detailed in the table 2.3 and in the following sub-sections.

Research	Expected signal	Energy region	Major backgrounds
Reactor antineutrino	60 IBDs/day	0–12 MeV	Radioactivity, cosmic muon
Supernova burst	5000 IBDs at 10 kpc	0–80 MeV	Negligible
DSNB (w/o PSD)	2300 elastic scattering		
Solar neutrino	2–4 IBDs/year	10–40 MeV	Atmospheric $\nu$
Atmospheric neutrino	hundreds per year for ${}^8\text{B}$	0–16 MeV	Radioactivity
Geoneutrino	hundreds per year	0.1–100 GeV	Negligible
	$\approx 400$ per year	0–3 MeV	Reactor $\nu$

TABLE 2.3 – Detectable neutrino signal in JUNO and the expected signal rates and major background sources

### Geoneutrinos

Geoneutrinos designate the antineutrinos coming from the decay of long-lived radioactive elements inside the Earth. The 1.8 MeV threshold necessary for the IBD makes it possible to measure geoneutrinos from  ${}^{238}\text{U}$  and  ${}^{232}\text{Th}$  decay chains. The studies of geoneutrinos can help refine the Earth crust models but is also necessary to characterise their signal, as they are a background to the mass ordering and oscillations parameters studies.

### Atmospheric neutrinos

Atmospheric neutrinos are neutrinos originating from the decay of  $\pi$  and  $K$  particles that are produced in extensive air showers initiated by the interactions of cosmic rays with the Earth atmosphere. Earth is mostly transparent to neutrinos below the PeV energy, thus JUNO will be able to see neutrinos coming from all directions. Their baseline range is large (15km  $\sim$  13000km), they can have energy between 0.1 GeV and 10 TeV and will contain all neutrino and antineutrinos flavour. Their studies is complementary to the reactor antineutrinos and can help refine the constraints on the NMO [2].

### Supernovae burst neutrinos

Neutrinos are crucial component during all stages of stellar collapse and explosion. Detection of neutrinos coming for core collapse supernovae will provide us important informations on the mech-

176 anisms at play in those events. Thanks to its 20 kt sensible volume, JUNO has excellent capabilities  
 177 to detect all flavour of the  $\mathcal{O}(10 \text{ MeV})$  postshock neutrinos, and using neutrinos of the  $\mathcal{O}(1 \text{ MeV})$   
 178 will give informations about the pre-supernovae neutrinos. All those informations will allow to  
 179 disentangle between the multiple hydro-dynamic models that are currently used to describe the  
 180 different stage of core-collapse supernovae.

### 181 Diffuse supernovae neutrinos background

182 Core-collapse supernovae in our galaxy are rare events, but they frequently occur throughout the  
 183 visible Universe sending burst of neutrinos in direction of the Earth. All those events contributes to  
 184 a low background flux of low-energy neutrinos called the Diffuse Supernovae Neutrino Background  
 185 (DSNB). Its flux and spectrum contains informations about the red-shift dependent supernovae rate,  
 186 the average supernovae neutrino energy and the fraction of black-hole formation in core-collapse su-  
 187 pernovae. Depending of the DSNB model, we can expect 2-4 IBD events per year in the energy range  
 188 above the reactor  $\bar{\nu}_e$  signal, which is competitive with the current Super-Kamiokande+Gadolinium  
 189 phase [17].

### 190 Beyond standard model neutrinos interactions

191 JUNO will also be able to probe for beyond standard model neutrinos interactions. After the main  
 192 physics topics have been accomplished, JUNO could be upgraded to probe for neutrinoless beta  
 193 decay ( $0\nu\beta\beta$ ). The detection of such event would give critical informations about the nature of  
 194 neutrinos, is it a majorana or a dirac particle. JUNO will also be able to probe for neutrinos that  
 195 would come for the decay or annihilation of Dark Matter inside the sun and neutrinos from putative  
 196 primordial black hole. Through the unitary test of the mixing matrix, JUNO will be able to search  
 197 for light sterile neutrinos. Thanks to JUNO sensitivity, multiple other exotic can be performed on  
 198 neutrino related beyond standard model interactions.

## 199 2.2 The JUNO detector

200 The JUNO detector is a scintillator detector buried 693.35 meters under the ground (1800 meters  
 201 water equivalent). It consist of Central Detector (CD), a water pool and a Top Tracker (TT) as showed  
 202 in figure 2.4a. The CD is an acrylic vessel containing the 20 ktons of Liquid Scintillator (LS). It is  
 203 supported by a stainless steel structure and is immersed in that water pool that is used as shielding  
 204 from external radiation and as a cherenkov detector for the background. The top of the experiment  
 205 is partially covered by the Top Tracker (TT), a plastic scintillator detector which is use to detect the  
 206 atmospheric muons background and is acting as a veto detector.

207 The top of the experiment also host the LS purification system, a water purification system, a ven-  
 208 tilation system to get rid of the potential radon in the air. The CD is observed by two system of  
 209 Photo-Multipliers Tubes (PMT). They are attached to the steel structure and their electronic readout  
 210 is submersed near them. A third system of PMT is also installed on the structure but are facing  
 211 outward of the CD, instrumenting the water to be cherenkov detector. The CD and the cherenkov  
 212 detector are optically separated by Tyvek sheet. A chimney for LS filling and purification and for  
 213 calibration operations connects the CD to the experimental hall from the top.

214 The CD has been dimensioned to meet the requirements presented in section 2.1.1:

- 215 — Its 20 ktons monolithic LS provide a volume sizeable enough, in combination with the ex-  
 216 pected  $\bar{\nu}_e$  flux, to reach the desired statistic in 6 years. Its monolithic nature also allow for a  
 217 full containment of most of the events, preventing the energy loss in non-instrumented parts  
 218 that would arise from a segmented detector.

- 219 — Its large overburden shield it from most of the atmospheric background that would pollute  
 220 the signal.  
 221 — The localization of the experiment, chosen to maximize the disappearance with a 53km base-  
 222 line and in a region that allow two nuclear power plant to be used as sources.

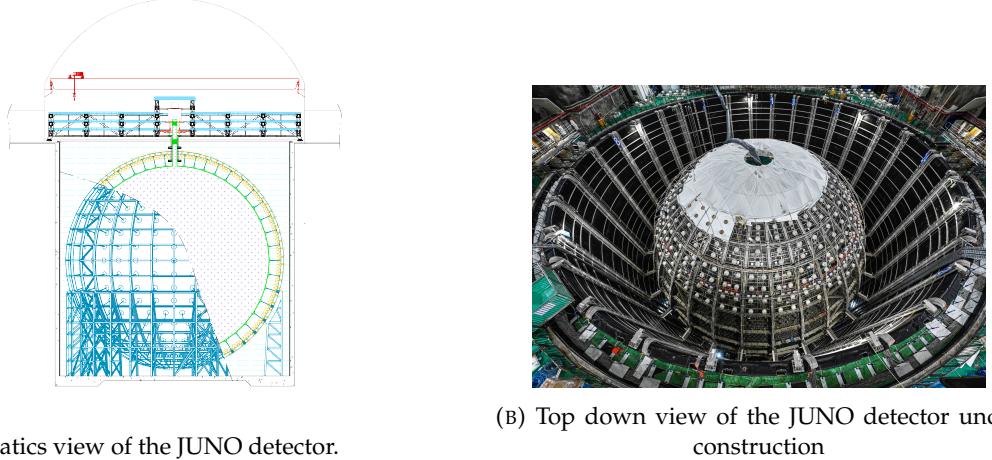


FIGURE 2.4

223 This section cover in details the different components of the detector and the detection systems.

### 224 2.2.1 Detection principle

The CD will detect the neutrino and measure their energy mainly via an Inverse Beta Decay (IBD) interaction with proton mainly from the  $^{12}\text{C}$  and H nucleus in the LS:

$$\bar{\nu}_e + p \rightarrow n + e^+$$

225 Kinematics calculation shows that this interaction has an energy threshold for the  $\bar{\nu}_e$  of  $(m_n + m_e -$   
 226  $m_p) \approx 1.806$  MeV [18] where  $m_\lambda$  is the mass of the  $\lambda$  particle. This threshold make the experiment  
 227 blind to very low energy neutrinos. The residual energy  $E_\nu - 1.806$  MeV is be distributed as kinetic  
 228 energy between the positron and the neutron. The energy of the emitted positron  $E_e$  is given by [18]

$$E_e = \frac{(E_\nu - \delta)(1 + \epsilon_\nu) + \epsilon_\nu \cos \theta \sqrt{(E_\nu - \delta)^2 + \kappa m_e^2}}{\kappa} \quad (2.2)$$

229 where  $\kappa = (1 + \epsilon_\nu)^2 - \epsilon_\nu^2 \cos^2 \theta \approx 1$ ,  $\epsilon_\nu = \frac{E_\nu}{m_p} \ll 1$  and  $\delta = \frac{m_n^2 - m_p^2 - m_e^2}{2m_p} \ll 1$ . We can see from this  
 230 equation that the positron energy is strongly correlated to the neutrino energy.

231 The positron and the neutron will then propagate in the detection medium, the Liquid Scintillator  
 232 (LS), loosing their kinetic energy by exciting the molecule of the LS (more details in section 2.2.2).  
 233 Once stopped, the positron will annihilate with an electron from the medium producing two 511  
 234 KeV gamma. Those gamma will themselves interact with the LS, exciting it before being absorbed  
 235 by photoelectrical effect. The neutron will be captured by an hydrogen, emitting a 2.2 MeV gamma  
 236 in the process. This gamma will also deposit its energy before being absorbed by the LS.

237 The scintillation photons have frequency in the UV and will propagate in the LS, being re-absorbed  
 238 and re-emitted by compton effect before finally be captured by PMTs instrumenting the acrylic  
 239 sphere. The analog signal of the PMTs digitized by the electronic is the signal of our experiment.

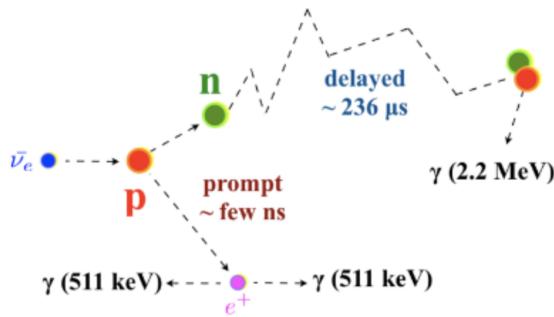


FIGURE 2.5 – Schematics of an IBD interaction in the central detector of JUNO

240 The signal produced by the positron is subsequently called the prompt signal, and the signal coming  
 241 from the neutron the delayed signal. This naming convention come from the fact that the positron  
 242 will deposit its energy rather quickly (few ns) where the neutron will take a bit more time ( $\sim 236 \mu\text{s}$ ).

### 243 2.2.2 Central Detector (CD)

244 The central detector, composed of 20 ktons of Liquid Scintillator (LS), is the main part of JUNO. The  
 245 LS is contained in a spherical acrylic vessel supported by a stainless steel structure. The CD and  
 246 its structural support are submerged in a cylindrical water pool of 43.5m diameter and 44m height.  
 247 We're confident that the water pool provide sufficient buffer protection in every direction against the  
 248 rock radioactivity.

#### 249 Acrylic vessel

250 The acrylic vessel is a spherical vessel of inner diameter of 35.4 m and a thickness of 120 mm. It is  
 251 assembled from 265 acrylic panels, thermo bonded together. The acrylic recipes has been carefully  
 252 tuned with extensive R&D to ensure it does not include plasticizer and anti-UV material that would  
 253 stop the scintillation photons. Those panels requires to be pure of radioactive materials to not  
 254 cause background. Current setup where the acrylic panels are molded in cleanrooms of class 10000,  
 255 let us reach a uranium and thorium contamination of <0.5 ppt. The molding and thermoforming  
 256 processes is optimized to increase the assemblage transparency in water to >96%. The acrylic vessel  
 257 is supported by a stainless steel structure via supporting node (fig 2.6). The structure and the nodes  
 258 are designed to be resilient to natural catastrophic events such as earthquake and can support many  
 259 times the effective load of the acrylic vessel.

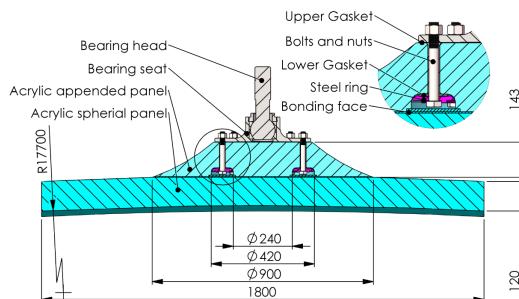


FIGURE 2.6 – Schematics of the supporting node for the acrylic vessel

260 **Liquid scintillator**

261 The Liquid Scintillator (LS) has a similar recipe as the one used in Daya Bay [19] but without gadolinium  
 262 doping. It is made of three components, necessary to shift the wavelength of emitted photons to  
 263 prevent their reabsorption:

- 264 1. The detection medium, the *linear alkylbenzene* (LAB). Selected because of its excellent trans-  
 265 parency, high flash point, low chemical reactivity and good light yield. Accounting for  $\sim$   
 266 98% of the LS, it is the main component with which ionizing particles and gamma interact.  
 267 Charged particles will collide with its electronic cloud transferring energy to the molecules,  
 268 gamma will interact via compton effect with the electronic cloud before finally be absorbed  
 269 via photoelectric effect.
- 270 2. The second component of the LS is the *2,5-diphenyloxazole* (PPO). A fraction of the excitation  
 271 energy of the LAB is transferred to the PPO, mainly via non radiative process [20]. The  
 272 PPO molecules de-excites in the same way, transferring their energy to the bis-MSB. The PPO  
 273 makes for 1.5 % of the LS.
- 274 3. The last component is the *p-bis(o-methylstyryl)-benzene* (bis-MSB). Once excited by the PPO, it  
 275 will emit photon with an average wavelength of  $\sim$  430 nm (full spectrum in figure 2.7) that  
 276 can be detected by our photo-multipliers systems. It amount for  $\sim$  0.5% of the LS.

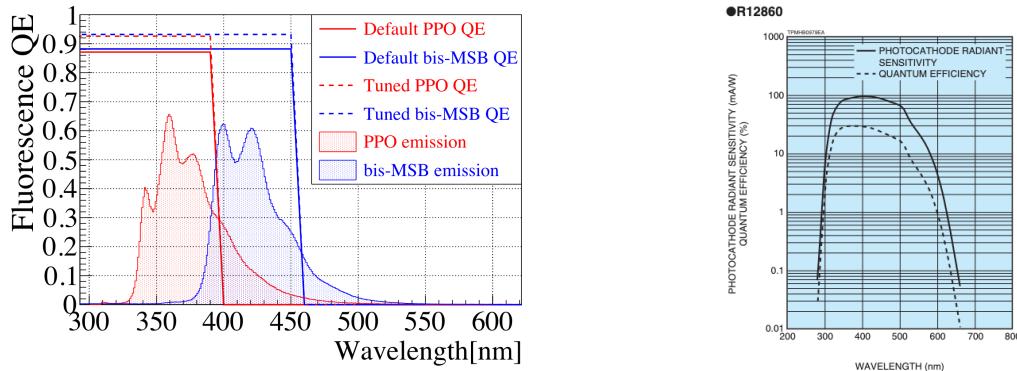


FIGURE 2.7 – On the left: Quantum efficiency (QE) and emission spectrum of the LAB and the bis-MSB [19]. On the right: Sensitivity of the Hamamatsu LPMT depending on the wavelength of the incident photons [21].

277 This formula has been optimized using dedicated studies with a Daya Bay detector [19, 22] to reach  
 278 the requirements for the JUNO experiment:

- 279 — A light yield / MeV of the amount of  $10^4$  photons to maximize the statistic in the energy  
 280 measurement.
- 281 — An attenuation length comparable to the size of the detector to prevent losing photons during  
 282 their propagation in the LS. The final attenuation length is 25.8m [23] to compare with the CD  
 283 diameter of 35.4m.
- 284 — Uranium/Thorium radiopurity to prevent background signal. The reactor neutrino program  
 285 require a contamination fraction  $F < 10^{-15}$  while the solar neutrino program require  $F <$   
 286  $10^{-17}$ .

287 The LS will frequently be purified and tested in the Online Scintillator Internal Radioactivity In-  
 288 vestigation System (OSIRIS) [24] to ensure that the requirements are kept during the lifetime of the  
 289 experiment, more details to be found in section 2.4.2.

290 **Large Photo-Multipliers Tubes (LPMTs)**

291 The scintillation light produced by the LS is then collected by Photo-Multipliers Tubes (PMT) that  
 292 transform the incoming photon into an electric signal. As described in figure 2.8, the incident photons  
 293 interact with the photocathode via photoelectric effect producing an electron called a Photo-Electron  
 294 (PE). This PE is then focused on the dynodes where the high voltage will allow it to be multiplied.  
 295 After multiple amplification the resulting charge - in coulomb [C] - is collected by the anode and  
 296 the resulting electric signal can be digitalized by the readout electronics from which the charge and  
 297 timing can be extracted.

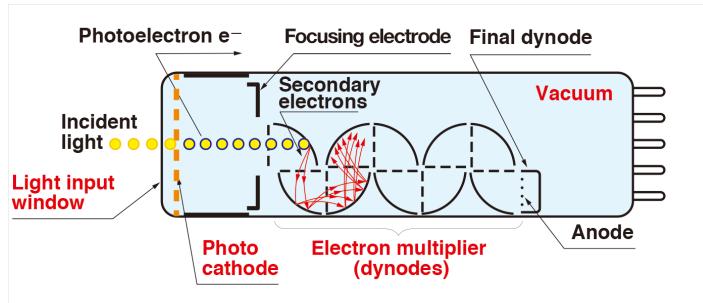


FIGURE 2.8 – Schematic of a PMT

298 The Large Photo-Multipliers Tubes (LPMT), used in the central detector and in the water pool, are  
 299 20-inch (50.8 cm) radius PMTs.  $\sim 5000$  dynode-PMTs [21] were produced by the Hamamatsu<sup>©</sup>  
 300 company and  $\sim 15000$  Micro-Channel Plate (MCP) [25] by the NNVT<sup>©</sup> company. This system is  
 301 the one responsible for the energy measurement with a energy resolution of  $3\%/\sqrt{E}$ , resolution  
 302 necessary for the mass ordering measurement. To reach this precision, the system is composed of  
 303 17612 PMTs quasi uniformly distributed over the detector for a coverage of 75.2% reaching  $\sim 1800$   
 304 PE/MeV or  $\sim 2.3\%$  resolution due to statistic, leaving  $\sim 0.7\%$  for the systematic uncertainties. They  
 305 are located outside the acrylic sphere in the water pool facing the center of the detector. To maintain  
 306 the resolution over the lifetime of the experiment, JUNO require a failure rate  $< 1\%$  over 6 years.

307 The LPMTs electronic are divided in two parts. One "near", located underwater, in proximity of the  
 308 LPMT to reduce the cable length between the PMT and early electronic. A second one, outside of the  
 309 detector that is responsible for higher level analysis before sending the data to the DAQ.

310 The light yield per MeV induce that a LPMT can collect between 1 and 1000 PE per event, causing  
 311 non linearity in the PMT response that need to be understood and calibrated, see section 2.3 for more  
 312 details.

313 **Small Photo-Multipliers Tubes (SPMTs)**

314 The Small PMT (SPMTs) system is made of 3-inch (7.62 cm) PMTs. They will be used in the CD  
 315 as a secondary detection system. Those 25600 SPMTs will observe the same events as the LPMTs,  
 316 thus sharing the physics and detector systematics up until the photon conversion. With a detector  
 317 coverage of 2.7%, this system will collect  $\sim 43$  PE/MeV for a final energy resolution of  $\sim 17\%$ .  
 318 This resolution is not enough to measure the NMO,  $\theta_{13}$ ,  $\Delta m_{31}^2$  but will be sufficient to independently  
 319 measure  $\theta_{12}$  and  $\Delta m_{21}^2$ .

320 Due to the low PE rate, SPMTs will be running in photo-counting mode in the reactor range and thus  
 321 will be insensitive to non-linearity effect. Using this property, the intrinsic charge non linearity of  
 322 the LPMTs can be measured by comparing the PE count in the SPMTs and LPMTs [26]. Also, due  
 323 to their smaller size and electronics, SPMTs have a better timing resolutions than the LPMTs. At

324 higher energy range, like supernovae events, LPMTs will saturate where SPMTs due to their lower  
325 PE collection will to produce a reliable measure of the energy spectrum.

326 The Data Acquisition System (DAQ) is designed to support the event rate of IBD, background, dark  
327 noise and supplementary storage buffers are present in the LPMT electronics to withstand the event  
328 rate during supernovae burst.

### 329 2.2.3 Veto detector

330 The CD will be bathed in constant background noise coming from numerous sources : the radioac-  
331 tivity from surrounding rock and its own components or from the flux of cosmic muons. This  
332 background needs to be rejected to ensure the purity of the IBD spectrum. To prevent a big part  
333 of them, JUNO use two veto detector that will tag events as background before CD analysis.

#### 334 Cherenkov in water pool

335 The Water Cherenkov Detector (WCD) is the instrumentation of the water buffer around the CD.  
336 When high speed charged particles will pass through the water, they will produced cherenkov  
337 photons. The light will be collected by 2400 MCP LPMTs installed on the outer surface of the CD  
338 structure. The muons veto strategy is based on a PMT multiplicity condition. WCD PMTs are  
339 grouped in ten zones: 5 in the top, 5 in the bottom. A veto is raised either when more than 19  
340 PMTs are triggered in one zone or when two adjacent zones simultaneously trigger more than 13  
341 PMTs. Using this trigger, we expect to reach a muon detection efficiency of 99.5% while keeping the  
342 noise at reasonable level.

#### 343 Top tracker

344 The JUNO Top Tracker (TT) is a plastic scintillator detector located on the top of the experiment (see  
345 figure 2.9). Made from plastic scintillator from OPERA [27] layered horizontally in 3 layers on the  
346 top of the detector, the TT will be able to detect incoming atmospheric muons. With its coverage,  
347 about 1/3 of the of all atmospheric muons that passing through the CD will also pass through the 3  
348 layer of the detector. While it does not cover the majority of the CD, the TT is particularly effective  
349 to detect muons coming through the filling chimney region which might present difficulties from the  
other subsystems in some classes of events.

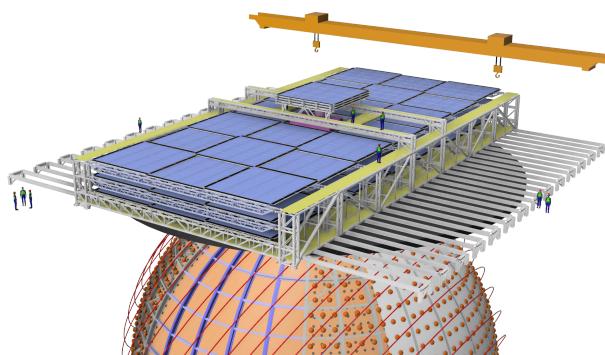


FIGURE 2.9 – The JUNO top tracker

### 351 2.3 Calibration strategy

352 The calibration is a crucial part of the JUNO experiment. Because we are looking at civil reactor  
 353 neutrino it might be impossible to run measurement without signal, it would need to shut down  
 354 every reactor from the Taishan and Yangjiang power plants which is realistically impossible. Because  
 355 of this continuous rate, low frequency signal event, we need high frequency, recognisable sources in  
 356 the energy range of interest : [0-12] MeV for the positron signal and 2.2 MeV for the neutron capture.  
 357 It is expected that the CD response will be different depending on the type of particle, due to the  
 358 interaction with LS, the position on the event and the optical response of the acrylic sphere (see  
 359 section 2.6). We also expect a non-linear energy response of the CD due to the LS properties [19] but  
 360 also due to the saturation of the LPMTs system when collecting a large amount of PE [26].

#### 361 2.3.1 Energy scale calibration

362 While electrons and positrons sources would be ideal, for a large LS detector thin-walled electrons  
 363 or positrons sources could lead to leakage of radionucleides causing radioactive contamination.  
 364 Instead, we consider gamma sources in the range of the prompt energy of IBDs. The sources are  
 365 reported in table 2.4.

Sources / Processes	Type	Radiation
$^{137}\text{Cs}$	$\gamma$	0.0662 MeV
$^{54}\text{Mn}$	$\gamma$	0.835 MeV
$^{60}\text{Co}$	$\gamma$	1.173 + 1.333 MeV
$^{40}\text{K}$	$\gamma$	1.461 MeV
$^{68}\text{Ge}$	$e^+$	annihilation 0.511 + 0.511 MeV
$^{241}\text{Am-Be}$	$n, \gamma$	neutron + 4.43 MeV ( $^{12}\text{C}^*$ )
$^{241}\text{Am-}^{13}\text{C}$	$n, \gamma$	neutron + 6.13 MeV ( $^{16}\text{O}^*$ )
$(n, \gamma)p$	$\gamma$	2.22 MeV
$(n, \gamma)^{12}\text{C}$	$\gamma$	4.94 MeV or 3.68 + 1.26 MeV

TABLE 2.4 – List of sources and their process considered for the energy scale calibration

366 For the  $^{68}\text{Ge}$  source, it will decay in  $^{68}\text{Ga}$  via electron capture, which will itself  $\beta^+$  decay into  $^{68}\text{Zn}$ .  
 367 The positrons will be absorbed by the enclosure so only the annihilation gamma will be released. In  
 368 addition,  $(\alpha, n)$  sources like  $^{241}\text{Am-Be}$  and  $^{241}\text{Am-}^{13}\text{C}$  are used to provide both high energy gamma  
 369 and neutrons, which will later be captured in the LS producing the 2.2 MeV gamma.

370 From this calibration we call  $E_{\text{vis}}$  the "visible energy" that is reconstructed by our current algorithms  
 371 and we compare it to the true energy deposited by the calibration source. The results shown in figure  
 372 2.10 show the expected response of the detector from calibration sources. The non-linearity is clearly  
 373 visible from the  $E_{\text{vis}}/E_{\text{true}}$  shape. See [28] for more details.

#### 374 2.3.2 Calibration system

375 The non-uniformity due to the event position in the detector (more details in section 2.6) will be  
 376 studied using multiples systems that are schematized in figure 2.11. They allow to position sources  
 377 at different location in the CD.

- 378 — For a one-dimension vertical calibration, the Automatic Calibration Unit (ACU) will be able  
 379 to deploy multiple radioactive sources or a pulse laser diffuser ball along the central axis of  
 380 the CD through the top chimney. The source position precision is less than 1cm.

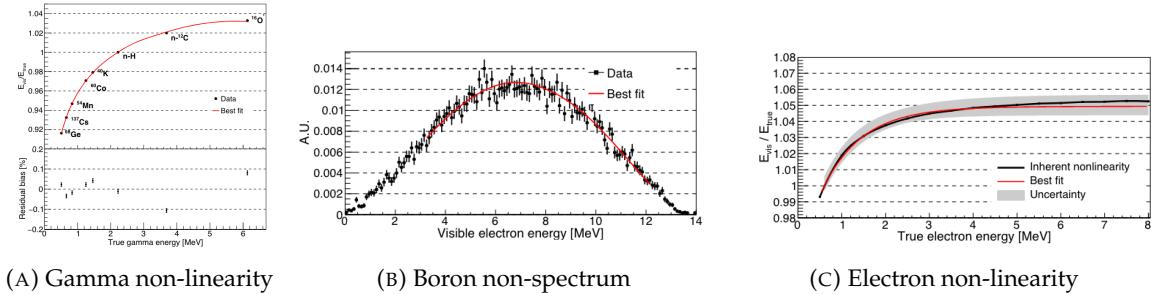


FIGURE 2.10 – Fitted and simulated non linearity of gamma, electron sources and from the  $^{12}\text{B}$  spectrum. Black points are simulated data. Red curves are the best fits

- For off-axis calibration, a calibration source attached to a Cable Loop System (CLS) can be moved on a vertical half-plane by adjusting the length of two connection cable. Two set of CSL will be deployed to provide a 79% effective coverage of a vertical plane.
- A Guiding Tube (GT) will surround the CD to calibrate the non-uniformity of the response at the edge of the detector
- A Remotely Operated under-LS Vehicle (ROV) can be deployed to desired location inside LS for a more precise and comprehensive calibration. The ROV will also be equipped with a camera for inspection of the CD.

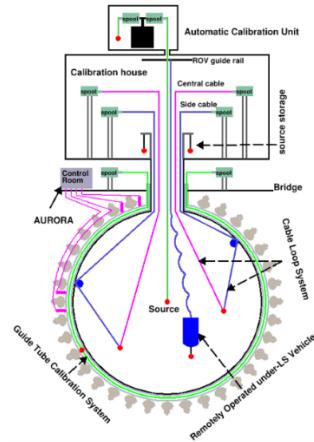


FIGURE 2.11 – Overview of the calibration system

- 389 The preliminary calibration program is depicted in table 2.5.

## 2.4 Satellite detectors

- 391 As introduced in section 2.1.1 and section 2.2.2, the precise knowledge and understanding of the  
392 detector condition is crucial for the measurements of the NMO and oscillation parameters. Thus two  
393 satellite detectors will be setup to monitor the experiment condition. TAO to monitor and understand  
394 the  $\bar{\nu}_e$  flux and spectrum coming from the nuclear reactor and OSIRIS to monitor the LS response.

Program	Purpose	System	Duration [min]
Weekly calibration	Neutron (Am-C)	ACU	63
	Laser	ACU	78
Monthly calibration	Neutron (Am-C)	ACU	120
	Laser	ACU	147
	Neutron (Am-C)	CLS	333
	Neutron (Am-C)	GT	73
Comprehensive calibration	Neutron (Am-C)	ACU, CLS and GT	1942
	Neutron (Am-Be)	ACU	75
	Laser	ACU	391
	$^{68}\text{Ge}$	ACU	75
	$^{137}\text{Cs}$	ACU	75
	$^{54}\text{Mn}$	ACU	75
	$^{60}\text{Co}$	ACU	75
	$^{40}\text{K}$	ACU	158

TABLE 2.5 – Calibration program of the JUNO experiment

### 395 2.4.1 TAO

396 The Taishan Antineutrino Observatory (TAO) [12, 29] is a ton-level gadolinium doped liquid scin-  
 397 tillator detector that will be located near the Taishan-1 reactor. It aim to measure the  $\bar{\nu}_e$  spectrum at  
 398 very low distance (< 30m) from the reactor to measure a quasi-unoscillated spectrum. TAO also aim  
 399 to provide a major contribution to the so-called reactor anomaly [13]. Its requirement are to the level  
 400 of 2 % energy resolution at 1 MeV.

401 **Detector**

402 The TAO detector is close, in concept, to the CD of JUNO. It is composed of an acrylic vessel  
 403 containing 2.8 tons of gadolinium-loaded LS instrumented by an array of silicon photomultipliers  
 404 (SiPM) reaching a 95% coverage. To efficiently reduce the dark count of those sensors, the detector  
 405 is cooled to -50 °C. The  $\bar{\nu}_e$  will interact with the LS via IBD, producing scintillation light, that will  
 406 be detected by the SiPMs. From this signal the  $\bar{\nu}_e$  energy and the full spectrum reconstructed. This  
 407 spectrum will then be used by JUNO to calibrate the unoscillated spectrum, most notably the fission  
 408 product fraction that impact the rate and shape of the spectrum. A schema of the detector is presented  
 409 in figure 2.12a.

410 **2.4.2 OSIRIS**

411 The Online Scintillator Internal Radioactivity Investigation System (OSIRIS) [24] is an ultralow back-  
 412 ground, 20 m<sup>3</sup> LS detector that will be located in JUNO cavern. It aim to monitor the radioactive  
 413 contamination, purity and overall response of the LS before it is injected in JUNO. OSIRIS will  
 414 be located at the end of the purification chain of JUNO, monitoring that the purified LS meet the  
 415 JUNO requirements. The setup is optimized to detect the fast coincidences decay of  $^{214}\text{Bi} - ^{214}\text{Po}$   
 416 and  $^{212}\text{Bi} - ^{212}\text{Po}$ , indicators of the decay chains of U and Th respectively.

417 **Detector**

418 OSIRIS is composed of an acrylic vessel that will contains 17t of LS. The LS is instrumented by  
 419 a PMT array of 64 20 inch PMTs on the top and the side of the vessel. To reach the necessary

background level required by the LS purity measurements, in addition to being 700m underground in the experiment cavern, the acrylic vessel is immersed in a tank of ultra pure water. The water is itself instrumented by another array of 20 inch PMTs, acting as muon veto. A schema of the detector is presented in figure 2.12b.

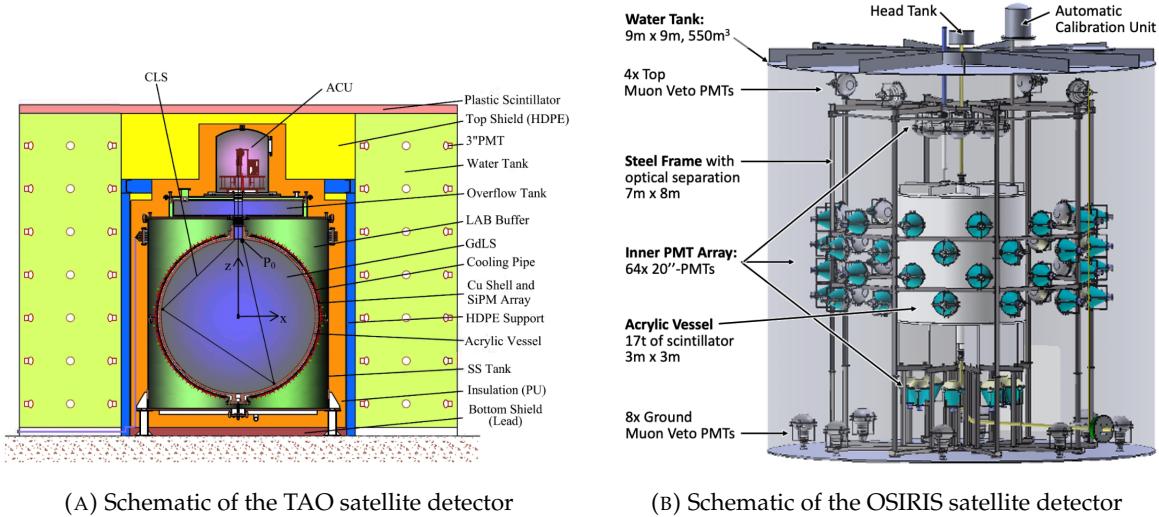


FIGURE 2.12

## 2.5 Software

The simulation, reconstruction and analysis algorithms are all packaged in the JUNO software, subsequently called the software. It is composed of multiple components integrated in the SNiPER [30] framework:

- Various primary particles simulators for the different kind of events, background and calibration sources.
- A Geant4 [31–33] Monte Carlo (MC) simulation containing the detectors geometries, a custom optical model for the LS and the supporting structures of the detectors. The Geant4 simulation integrate all relevant physics process for JUNO, validated by the collaboration. This step of the simulation is commonly called *Detsim* and compute up to the production of photo-electrons in the PMTs. The optics properties of the different materials and detector components have been measured beforehand to be used to define the material and surfaces in the simulation.
- An electronic simulation, simulating the response waveform of the PMTs, tracking it through the digitization process, accounting for effects such as non-linearity, dark noise, Time Transit Spread (TTS), pre-pulsing, after-pulsing and ringing if the waveform. It's also the step handling the event triggers and mixing. This step is commonly referenced as *Elecsim*.
- A waveform reconstruction where the digitized waveform are filtered to remove high-frequency white noise and then deconvoluted to yield time and charge informations of the photons hits on the PMTs. This step is commonly referenced as *Calib*.
- The charge and time informations are used by reconstruction algorithms to reconstruct the interaction vertex and the deposited energy. This step is commonly reported as *Reco*. See section 2.6 for more details on the reconstruction.
- Once the singular events are reconstructed, they go through event pairing and classification to select IBD events. This step is named Event Classification.

448 — The purified signal is then analysed by the analysis framework which depend of the physics  
 449 topic of interest.

450 The steps Reco and Event Classification are divided into two category of algorithm. Fast but less  
 451 accurate algorithms that are running during the data taking designated as the *Online* algorithms.  
 452 Those algorithm are used to take the decision to save the event on tape or to throw it away. More  
 453 accurate algorithms that run on batch of events designated *Offline* algorithms. They are used for the  
 454 physics analysis. The Offline Reco will be one of the main topic of interest for this thesis.

## 455 2.6 State of the art of the Offline IBD reconstruction in JUNO

456 The main reconstruction method currently run in JUNO is a data-driven method based on a like-  
 457 lihood maximization [34, 35] using only the LPMTs. The first step is to reconstruct the interaction  
 458 vertex from which the energy reconstruction is dependent. It is also necessary for event pairing and  
 459 classification.

### 460 2.6.1 Interaction vertex reconstruction

461 To start the likelihood maximization, a rough estimation of the vertex and of the event timing is  
 462 needed. We start by estimating the vertex position using a charge based algorithm.

#### 463 Charge based algorithm

464 The charge-based algorithm is basically base on the charge-weighted average of the PMT position.

$$\vec{r}_{cb} = a \cdot \frac{\sum_i q_i \cdot \vec{r}_i}{\sum_i q_i} \quad (2.3)$$

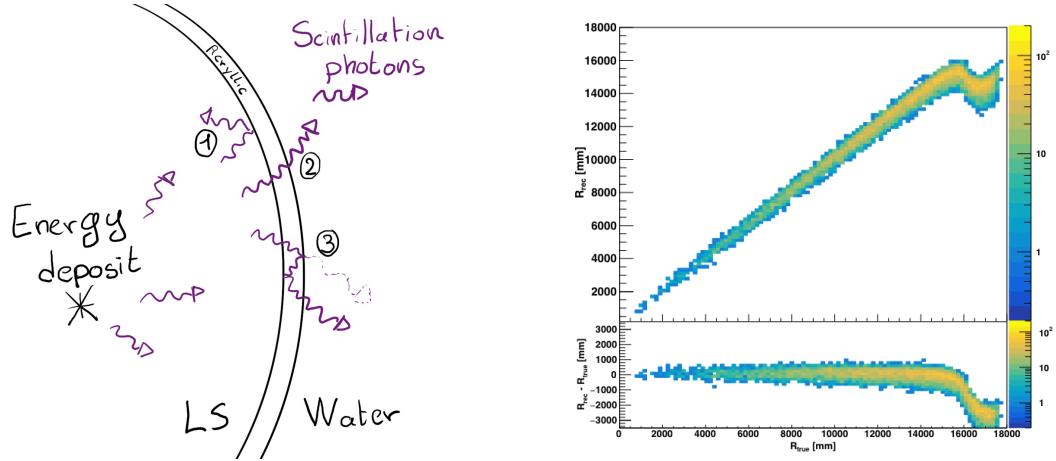
465 Where  $q_i$  is the reconstructed charge of the pulse of the  $i$ th PMT and  $\vec{r}_i$  is its position.  $\vec{r}_0$  is the  
 466 reconstructed interaction position.  $a$  is a scale factor introduced because a weighted average over  
 467 a 3D sphere is inherently biased. Using calibration we can estimate  $a \approx 1.3$  [36]. The results in  
 468 figure 2.13b shows that the reconstruction is biased from around 15m and further. This is due to the  
 469 phenomena called “total reflection area” or TR Area.

470 As depicted in the figure 2.13a the optical photons, given that they have a sufficiently large incidence  
 471 angle, can be deviated of their trajectories when passing through the interfaces LS-acrylic and water-  
 472 acrylic due to the optical index difference. This cause photons to be lost or to be detected by PMT  
 473 further than anticipated if we consider their rectilinear trajectories. This cause the charge barycenter  
 474 the be located closer to the center than the event really is.

475 It is to be noted that charge based algorithm, in addition to be biased near the edge of the detector,  
 476 does not provide any information about the timing of the event. Therefore, a time based algorithm  
 477 needs to be introduced to provide initial values.

#### 478 Time based algorithm

479 The time based algorithm use the distribution of the time of flight corrections  $\Delta t$  (Eq 2.4) of an event  
 480 to reconstruct its vertex and  $t_0$ . It follow the following iterations:



(A) Illustration of the different optical photons reflection scenarios. 1 is the reflection of the photon at the interface LS-acrylic or acrylic-water. 2 is the transmission of the photons through the interfaces. 3 is the conduction of the photon in the acrylic.

(B) Heatmap of  $R_{rec}$  and  $R_{rec} - R_{true}$  as a function of  $R_{true}$  for 4MeV prompt signals uniformly distributed in the detector calculated by the charge based algorithm

FIGURE 2.13

481 1. Use the charge based algorithm to get an initial vertex to start the iteration.

482 2. Calculate the time of flight correction for the  $i$ th PMT using

$$\Delta t_i(j) = t_i - \text{tof}_i(j) \quad (2.4)$$

483 where  $j$  is the iteration step,  $t_i$  is the timing of the  $i$ th PMT, and  $\text{tof}_i$  is the time-of-flight of the  
484 photon considering an rectilinear trajectory and an effective velocity in the LS and water (see  
485 [36] for detailed description of this effective velocity). Plot the  $\Delta t$  distribution and label the  
486 peak position as  $\Delta t^{\text{peak}}$  (see fig 2.14a).

487 3. Calculate a correction vector  $\vec{\delta}[\vec{r}(j)]$  as

$$\vec{\delta}[\vec{r}(j)] = \frac{\sum_i \left( \frac{\Delta t(j) - \Delta t^{\text{peak}}(j)}{\text{tof}_i(j)} \right) \cdot (\vec{r}_0(j) - \vec{r}_i)}{N^{\text{peak}}(j)} \quad (2.5)$$

488 where  $\vec{r}_0$  is the vertex position at the beginning of this iteration,  $\vec{r}_i$  is the position of the  $i$ th  
489 PMT. To minimize the effect of scattering, dark noise and reflection, only the pulse happening  
490 in a time window (-10 ns, +5 ns) around  $\Delta t^{\text{peak}}$  are considered.  $N^{\text{peak}}$  is the number of PE  
491 collected in this time-window.

492 4. if  $\vec{\delta}[\vec{r}(j)] < 1\text{mm}$  or  $j \geq 100$ , stop the iteration. Otherwise  $\vec{r}_0(j+1) = \vec{r}_0(j) + \vec{\delta}[\vec{r}(j)]$  and go to  
493 step 2.

494 However because the earliest arrival time is used,  $t_i$  is related to the number photoelectrons  $N_i^{\text{pe}}$   
495 detected by the PMT [37–39]. To reduce bias in the vertex reconstruction, the following equation is  
496 used to correct  $t_i$  into  $t'_i$ :

$$t'_i = t_i - p_0 / \sqrt{N_i^{\text{pe}}} - p_1 - p_2 / N_i^{\text{pe}} \quad (2.6)$$

497 The parameters  $(p_0, p_1, p_2)$  were optimized to (9.42, 0.74, -4.60) for Hamamatsu PMTs and (41.31,  
498 -12.04, -20.02) for NNVT PMTs [36]. The results presented in figure 2.14b shows that the time based

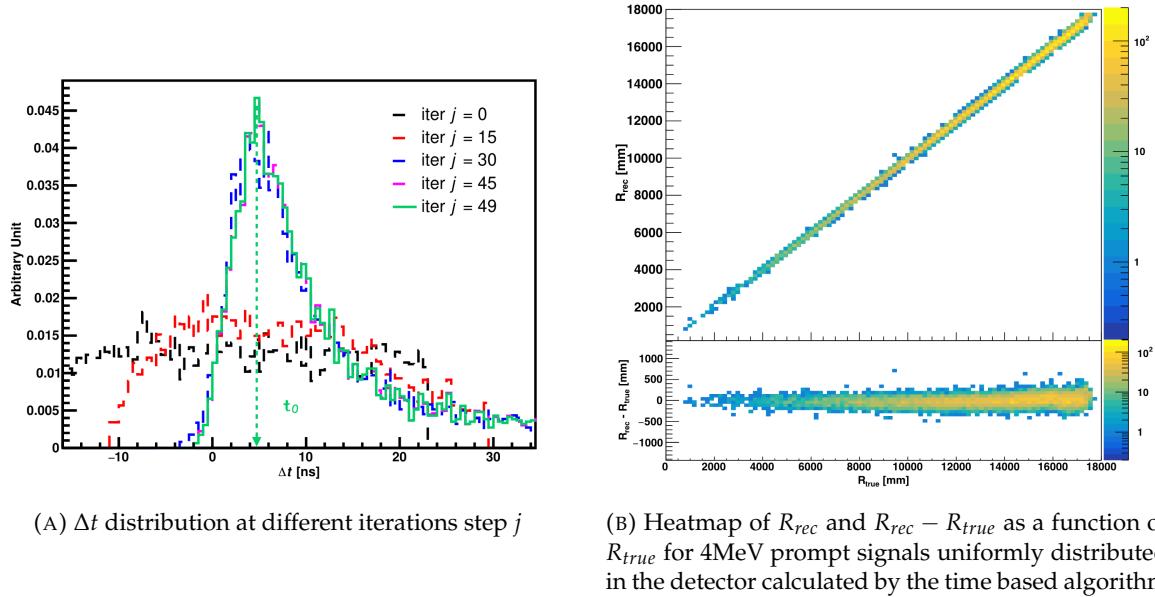
(A)  $\Delta t$  distribution at different iterations step  $j$ (B) Heatmap of  $R_{rec}$  and  $R_{rec} - R_{true}$  as a function of  $R_{true}$  for 4MeV prompt signals uniformly distributed in the detector calculated by the time based algorithm

FIGURE 2.14

499 algorithm provide a more accurate vertex and is unbiased even in the TR area. This results  $(\vec{r}_0, t_0)$  is  
500 used as initial value for the likelihood algorithm.

### 501 Time likelihood algorithm

502 The time likelihood algorithm use the residual time expressed as follow

$$t_{res}^i(\vec{r}_0, t_0) = t_i - \text{tof}_i - t_0 \quad (2.7)$$

503 In a first order approximation, the scintillator time response Probability Density Function (PDF) can  
504 be described as the emission time profile of the scintillation photons, the Time Transit Spread (TTS)  
505 and the dark noise of the PMTs. The emission time profile  $f(t_{res})$  is described like

$$f(t_{res}) = \sum_k \frac{\rho_k}{\tau_k} e^{-\frac{t_{res}}{\tau_k}}, \sum_k \rho_k = 1 \quad (2.8)$$

506 as the sum of the  $k$  component that emit light in the LS each one characterised by it's decay time  $\tau_k$   
507 and intensity fraction  $\rho_k$ . The TTS component is expressed as a gaussian convolution

$$g(t_{res}) = \frac{1}{\sqrt{2\pi}\sigma} e^{-\frac{(t_{res}-\nu)^2}{2\sigma^2}} \cdot f(t_{res}) \quad (2.9)$$

508 where  $\sigma$  is the TTS of PMTs and  $\nu$  is the average transit time. The dark noise is not correlated with any  
509 physical events and considered as constant rate over the time window considered  $T$ . By normalizing  
510 the dark noise probability  $\epsilon(t_{res})$  as  $\int_T \epsilon(t_{res}) dt_{res} = \epsilon_{dn}$ , it can be integrated in the PDF as

$$p(t_{res}) = (1 - \epsilon_{dn}) \cdot g(t_{res}) + \epsilon(t_{res}) \quad (2.10)$$

511 The distribution of the residual time  $t_{res}$  of an event can then be compared to  $p(t_{res})$  and the best

512 fitting vertex  $\vec{r}_0$  and  $t_0$  can be chosen by minimizing

$$\mathcal{L}(\vec{r}_0, t_0) = -\ln \left( \prod_i p(t_{\text{res}}^i) \right) \quad (2.11)$$

513 The parameter of Eq. 2.10 can be measured experimentally. The results shown in figure 2.15 used  
 514 PDF from monte carlo simulation. The results shows that  $R_{\text{rec}} - R_{\text{true}}$  is biased depending on the  
 515 energy. While this could be corrected using calibration, another algorithm based on charge likelihood  
 516 was developed to correct this problem.

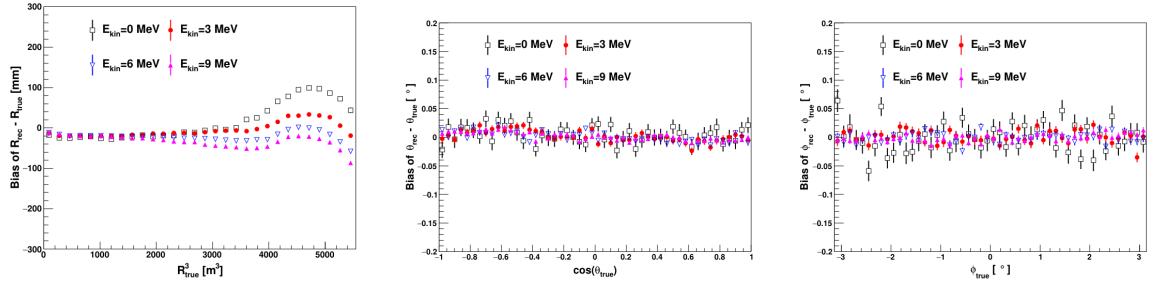


FIGURE 2.15 – Bias of the reconstructed radius  $R$  (left),  $\theta$  (middle) and  $\phi$  (right) for multiple energies by the time likelihood algorithm

### 517 Charge likelihood algorithm

518 Similarly to the time likelihood algorithms that use a timing PDF, the charge likelihood algorithm  
 519 use a PE PDF for each PMT depending on the energy and position of the event. With  $\mu(\vec{r}_0, E)$  the  
 520 mean expected number of PE detected by each PMT, the probability to observe  $N_{pe}$  in a PMT follow  
 521 a Poisson distribution. Thus

- 522 — The probability to observe no hit ( $N_{pe} = 0$ ) in the  $j$ th PMT is  $P_{\text{nohit}}^j(\vec{r}_0, E) = e^{-\mu_j}$
- 523 — The probability to observe  $N_{pe} \neq 0$  in the  $i$ th PMT is  $P_{\text{hit}}^i(\vec{r}_0, E) = \frac{\mu^{N_{pe}} e^{-\mu_i}}{N_{pe}!}$

524 Therefore, the probability to observe a specific hit pattern can be expressed as

$$P(\vec{r}_0, E) = \prod_j P_{\text{nohit}}^j(\vec{r}_0, E) \cdot \prod_i P_{\text{hit}}^i(\vec{r}_0, E) \quad (2.12)$$

525 The best fit values of  $\vec{R}_0$  and  $E$  can then be calculated by minimizing the negative log-likelihood

$$\mathcal{L}(\vec{r}_0, E) = -\ln(P(\vec{r}_0, E)) \quad (2.13)$$

526 In principle,  $\mu_i(\vec{r}_0, E)$  could be expressed

$$\mu_i(\vec{r}_0, E) = Y \cdot \frac{\Omega(\vec{r}_0, r_i)}{4\pi} \cdot \epsilon_i \cdot f(\theta_i) \cdot e^{-\sum_m \frac{d_m}{\zeta_m}} \cdot E + \delta_i \quad (2.14)$$

527 where  $Y$  is the energy scale factor,  $\Omega(\vec{r}_0, r_i)$  is the solid angle of the  $i$ th PMT,  $\epsilon_i$  is its detection  
 528 efficiency,  $f(\theta_i)$  its angular response,  $\zeta_m$  is the attenuation length in the materials and  $\delta_i$  the expected  
 529 number of dark noise.

530 However Eq. 2.14 assume that the scintillation light yield is linear with energy and describe poorly  
 531 the contribution of indirect light, shadow effect due to the supporting structure and the total reflec-

tion effects. The solution is to use data driven methods to produce the pdf by using the calibrations sources and position described in section 2.3. In the results presented in figures 2.16, the PDF was produced using MC simulation and 29 specific calibrations position [36] along the Z-axis of the detector. We see that the charge likelihood algorithm show little bias in the TR area and a better

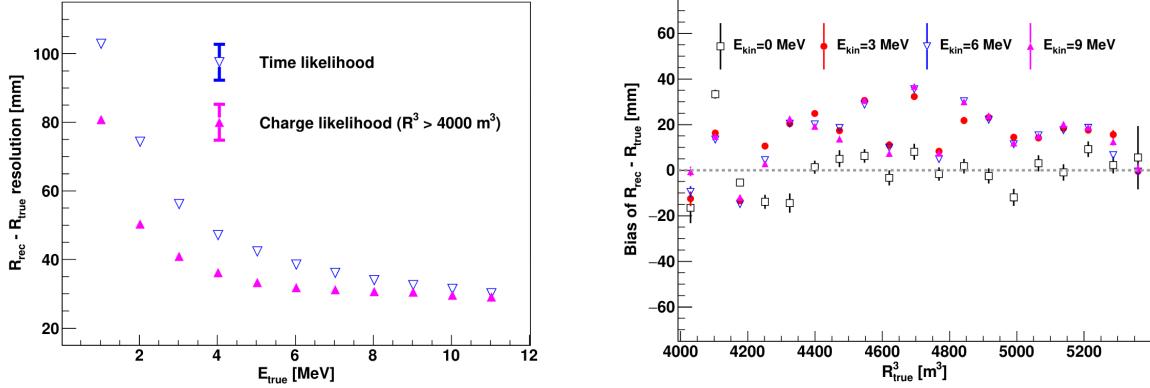


FIGURE 2.16 – On the left: Resolution of the reconstructed  $R$  as a function of the energy in the TR area ( $R^3 > 4000 \text{ m}^3 \equiv R > 16\text{m}$ ) by the charge and time likelihood algorithms. On the right: Bias of the reconstructed  $R$  in the TR area for different energies by the charge likelihood algorithm

resolution than the time likelihood. The figure 2.17 shows the radial resolution of the different algorithm presented for this section, we can see the refinement at each step and that the charge likelihood yield the best results.

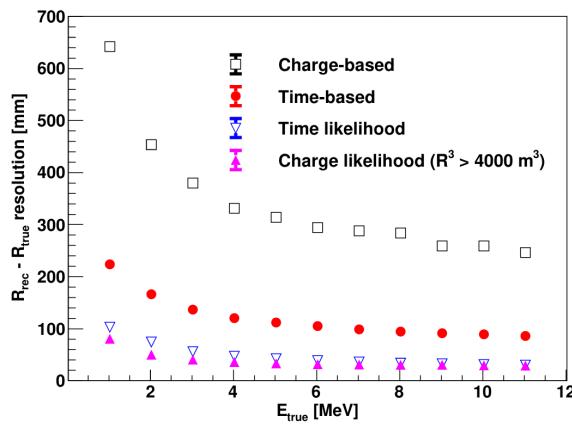


FIGURE 2.17 – Radial resolution of the different vertex reconstruction algorithms as a function of the energy

The charge based likelihood algorithms already give use some information on the energy as Eq. 2.13 is minimized but the energy can be further refined as shown in the next section.

## 2.6.2 Energy reconstruction

As explained in section 2.1.1, energy resolution is crucial for the NMO and oscillation parameters measurements. Thus the energy reconstruction algorithm should take into consideration as much

544 detector effect as possible. The following method is a data driven method based on calibration  
 545 samples inspired by the charge likelihood algorithm described above [40].

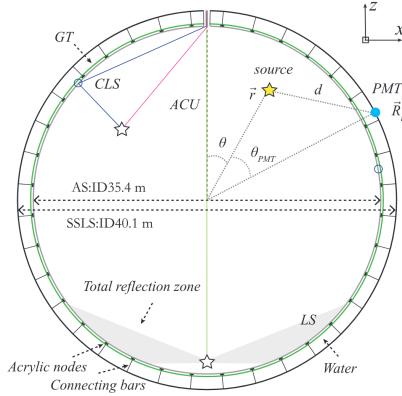


FIGURE 2.18 – Definition of the variables used in the energy reconstruction

### 546 Charge estimation

547 The most important element in the energy reconstruction is  $\mu_i(\vec{r}_0, E)$  described in Eq. 2.14. For  
 548 realistic cases, we also need to take into account the electronics effect that were omitted in the  
 549 previous section. Those effect will cause a charge smearing due to the uncertainties in the  $N_{pe}$   
 550 reconstruction. Thus we define  $\hat{\mu}^L(\vec{r}_0, E)$  which is the expected  $N_{pe}/E$  in the whole detector for an  
 551 event with visible energy  $E_{vis}$  and position  $\vec{r}_0$ . The position of the event and PMTs are now defined  
 552 using  $(r, \theta, \theta_{pmt})$  as defined in figure 2.18.

$$\hat{\mu}(r, \theta, \theta_{pmt}, E_{vis}) = \frac{1}{E_{vis}} \frac{1}{M} \sum_i^M \frac{\bar{Q}_i - \mu_i^D}{DE_i}, \quad \mu_i^D = DNR_i \cdot L \quad (2.15)$$

553 where  $i$  runs over the PMTs with the same  $\theta_{pmt}$ ,  $DE_i$  is the detection efficiency of the  $i$ th PMT.  $\mu_i^D$   
 554 is the expected number of dark noise photoelectrons in the time window  $L$ . The time window have  
 555 been optimized to  $L = 280$  ns [40].  $\bar{Q}_i$  is the average recorded photoelectrons in the time window  
 556 and  $\hat{Q}_i$  is the expected average charge for 1 photoelectron. The  $N_{pe}$  map is constructed following the  
 557 procedure described in [35].

### 558 Time estimation

559 The second important observable is the hit time of photons that was previously defined in Eq. 2.7. It  
 560 is here refined as

$$t_r = t_h - \text{tof} - t_0 = t_{LS} + t_{TT} \quad (2.16)$$

561 where  $t_h$  is the time of hit,  $t_{LS}$  is the scintillation time and  $t_{TT}$  the transit time of PMTs that is described  
 562 by a gaussian

$$t_{TT} = \mathcal{N}(\mu_{TT} + t_d, \sigma_{TT}) \quad (2.17)$$

563 where  $\mu_{TT}$  is the mean transit time in PMTs,  $\sigma_{TT}$  is the Transit Time Spread (TTS) of the PMTs and  $t_d$   
 564 is the delay time in the electronics. The effective refraction index of the LS is also corrected to take  
 565 into account the propagation distance in the detector.

566 The timing PDF  $P_T(t_r|r, d, \mu_l, \mu_d, k)$  can now be generated using calibration sources [40]. This PDF  
 567 describe the probability that the residual time of the first photon hit is in  $[t_r, t_r + \delta]$  with  $r$  the radius  
 568 of the event vertex,  $d = |\vec{r} - \vec{r}_{PMT}|$  the propagation distance,  $\mu_l$  and  $\mu_d$  the expected number of PE  
 569 and dark noise in the electronic reading window and  $k$  is the detected number of PE.

570 Now let denote  $f(t, r, d)$  the probability density function of "photoelectron hit a time  $t$ " for an event  
 571 happening at  $r$  where the photons traveled the distance  $d$  in the LS

$$F(t, r, d) = \int_t^L f(t', r, d) dt' \quad (2.18)$$

572 Based on the PDF for one photon  $k = 1$ , one can define

$$P_T^l(t|k = n) = I_n^l [f_l(t) F_l^{n-1}(t)] \quad (2.19)$$

573 where the indicator  $l$  means that the photons comes from the LS and  $I_n^l$  a normalisation factor. To this  
 574 pdf we add the probability to have photons coming from the dark noise indicated by the indicator  $d$   
 575 using

$$f_d(t) = 1/L, F_d(t) = 1 - \frac{t}{L} \quad (2.20)$$

576 and so for the case where only one photon is detected by the PMT ( $k = 1$ )

$$P_T(t|\mu_l, \mu_d, k = 1) = I_1[P(1, \mu_l)P(0, \mu_d)f_l(t) + P(0, \mu_l)P(1, \mu_d)f_d(t)] \quad (2.21)$$

577 where  $P(k_\alpha, \mu_\alpha)$  is the Poisson probability to detect  $k_\alpha$  PE from  $\alpha \in \{l, d\}$  with the condition  $k_l + k_d =$   
 578  $k$ .

579 Now that we have the individual timing and charge probability we can construct the charge likeli-  
 580 hood referred as QMLE:

$$\mathcal{L}(q_1, q_2, \dots, q_N | \vec{r}, E_{vis}) = \prod_{j \in \text{unfired}} e^{-\mu_j} \prod_{i \in \text{fired}} \left( \sum_{k=1}^K P_Q(q_i|k) \cdot P(k, \mu_i) \right) \quad (2.22)$$

581 where  $\mu_i = E_{vis}\hat{\mu}_i^L + \mu_i^D$  and  $P(k, \mu_i)$  is the Poisson probability of observing  $k$  PE.  $P_Q(q_i|k)$  is the  
 582 charge pdf for  $k$  PE. And we can also construct the time likelihood referred as TMLE:

$$\mathcal{L}(t_{1,r}, t_{2,r}, \dots, t_{N,r} | \vec{r}, t_0) = \prod_{i \in \text{hit}} \frac{\sum_{k=1}^K P_T(t_{i,r}|r, d, \mu_i^l, \mu_i^d, k) \cdot P(k, \mu_i^l + \mu_i^d)}{\sum_{k=1}^K P(k, \mu_i^l + \mu_i^d)} \quad (2.23)$$

583 where  $K$  is cut to 20 PE and hit is the set of hits satisfying  $-100 < t_{i,r} < 500$  ns.

584 Merging those two likelihood give the charge-time likelihood QTMLLE

$$\mathcal{L}(q_1, q_2, \dots, q_N; t_{1,r}, t_{2,r}, \dots, t_{N,r} | \vec{r}, t_0, E_{vis}) = \mathcal{L}(q_1, q_2, \dots, q_N | \vec{r}, E_{vis}) \cdot \mathcal{L}(t_{1,r}, t_{2,r}, \dots, t_{N,r} | \vec{r}, t_0) \quad (2.24)$$

585 The radial and energy resolutions of the different likelihood are presented in figure 2.19 (from [40]).  
 586 We can see the improvement of adding the time information to the vertex reconstruction and that  
 587 an increase in vertex precision can bring improvement in the energy resolution, especially at low  
 588 energies.

589 Data driven methods prove to be performant in the energy and vertex reconstruction given that we  
 590 have enough calibrations sources to produce the PDF. In the next section, we'll see another type of  
 591 data-driven method based on machine learning.

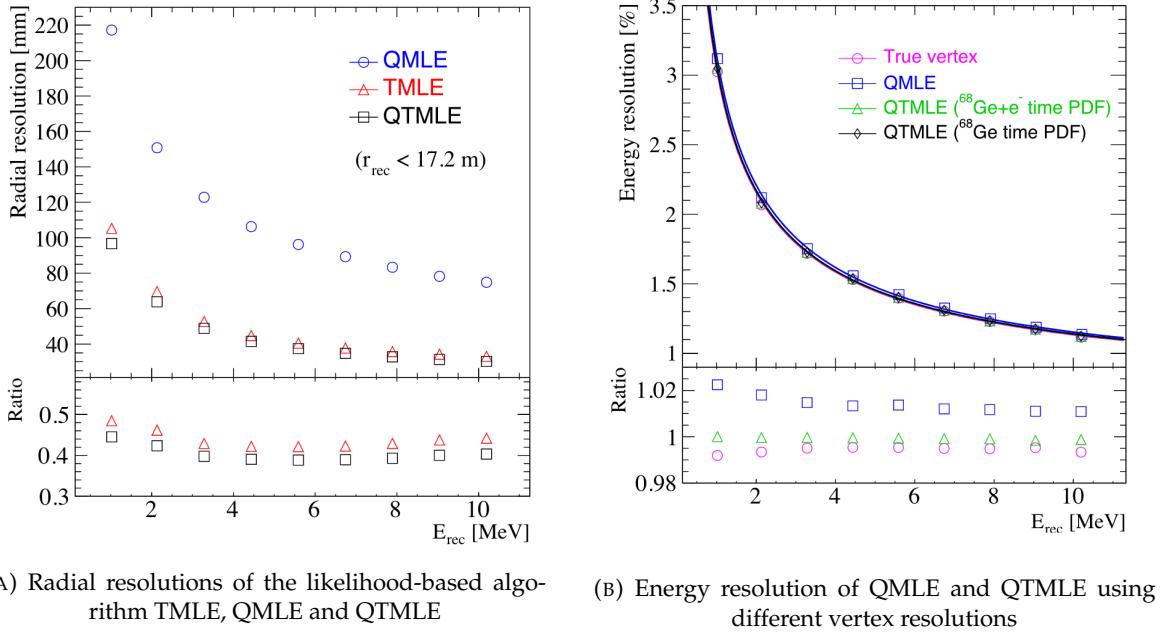


FIGURE 2.19

### 2.6.3 Machine learning for reconstruction

Machine learning (ML) is a family of data-driven algorithms that are inferring behavior and results from a training dataset. A overview of methods and detailed explanation of the Neural Network (NN) subfamily can be found in Chapter 3.

The power of ML is the ability to model complex response to a specific problem. In JUNO the reconstruction problematic can be expressed as follow: knowing that each PMT, large or small, detected a given number of PE  $Q$  at a given time  $t$  and their position is  $x, y, z$  where did the energy was deposited and how much energy was it, modeling a function that naively goes:

$$\mathbb{R}^{5 \times N_{\text{pmt}}} \mapsto \mathbb{R}^4 \quad (2.25)$$

It is worth pointing that while this is already a lot in informations, this is not the rawest representation of the experiment. We could indeed replace the charge and time by the waveform in the time window of the event but that would lead to an input representation size that would exceed our computational limits. Also, due to those computational limits, most of the ML algorithm reduce this input phase space either by structurally encoding the information (pictures, graph), by aggregating it (mean, variance, ...) or by exploiting invariance and equivariance of the experiment (rotational invariance due to the sphericity, ...).

For machine learning to converge to performant algorithm, a large dataset exploring all the phase space of interest is needed. For the following studies, data from the monte carlo simulation presented in section 2.5 are used for training. When the detector will be finished calibrations sources will be complementarily be used.

611 **Boosted Decision Tree (BDT)**

612 On of the most classic ML method used in physics in last years is the Boosted Decision Tree (see  
 613 chapter 3.1). They have been explored for vertex reconstruction [41] et for energy reconstruction [41,  
 614 42].

615 For vertex and energy reconstruction a BDT was developed using the aggregated informations pre-  
 616 sented in 2.6.

Parameter	description
$nHits$	Total number of hits
$x_{cc}, y_{cc}, z_{cc}, R_{cc}$	Coordinates of the center of charge
$ht_{mean}, ht_{std}$	Hit time mean and standard deviation

TABLE 2.6 – Features used by the BDT for vertex reconstruction

617 Its reconstruction performances are presented in figure 2.21.

618 A second and more advanced BDT, subsequently named BDTE, that only reconstruct energy use a  
 619 different set of features [42]. They are presented in the table 2.7

620 **Neural Network (NN)**

621 The physics have shown a rising for Neural Network (NN) in the past years for event reconstruction,  
 622 notably in the neutrino community [43–46]. Three type of neural networks have explored for event  
 623 reconstruction in JUNO Deep Neural Network (DNN), Convolutional Neural Network (CNN) and  
 624 Graph Network (GNN). More explanation about those neural network can be found in chapter 3.

625 The CNN are using 2D projection of the detector representing it as an image with two channel, one  
 626 for the charge  $Q$  and one for the time  $t$ . The position of the PMTs is structurally encoded in the pixel  
 627 containing the information of this PMT. In [41], the pixel is chosen based on a transformation of  $\theta$   
 628 and  $\phi$  coordinates to the 2D plane and rounded to the nearest pixel. A sufficiently large image has  
 629 been chosen to prevent two PMT to be located in the same pixel. An example of this projection can  
 630 be found in figure 2.20. The performances of the CNN can be found in figure 2.21.

631 Using 2D have the upside of encoding a large part of the informations structurally but loose the rota-  
 632 tional invariance of the detector. It also give undefined information to the neural network (what is a  
 633 pixel without PMT ? What should be its charge and time ?), cause deformation in the representation  
 634 of the detector (sides of projection) and loose topological informations.

635 One of the way to present structurally the sphericity of JUNO to a NN is to use a graph: A collection  
 636 of objects  $V$  called nodes and relations  $E$  called edges, each relation associated to a couple  $v_1, v_2$

AccumCharge	$ht_{5\%-2\%}$
$R_{cht}$	$pe_{mean}$
$z_{cc}$	$J_{cht}$
$pe_{std}$	$\phi_{cc}$
nPMTs	$ht_{35\%-30\%}$
$ht_{kurtosis}$	$ht_{20\%-15\%}$
$ht_{25\%-20\%}$	$pe_{35\%}$
$R_{cc}$	$ht_{30\%-25\%}$

TABLE 2.7 – Features used by the BDTE algorithm.  $pe$  and  $ht$  reference the charge  
 and hit-time distribution respectively and the percentages are the quantiles of those  
 distributions.  $cht$  and  $cc$  reference the barycenters of hit time and charge respectively

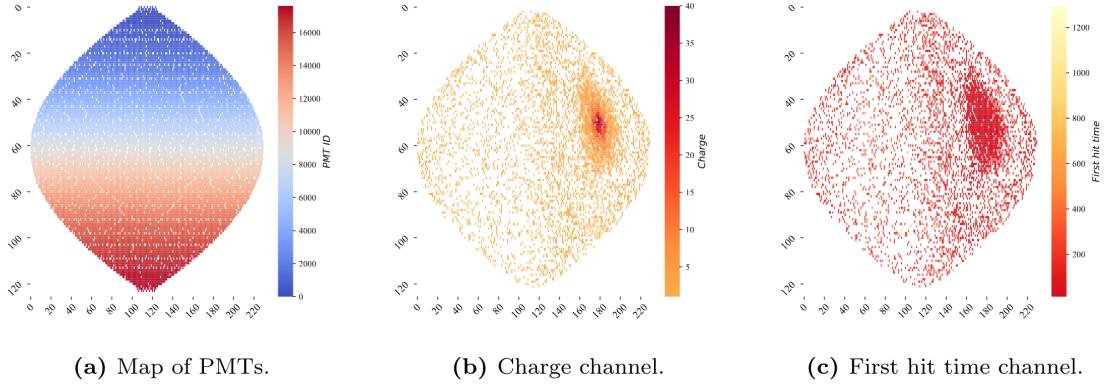


FIGURE 2.20 – Projection of the LPMTs in JUNO on a 2D plane. (a) Show the distribution of all PMTs and (b) and (c) are example of what the charge and time channel looks like respectively

637 forming the graph  $G(E, V)$ . Nodes and edges can hold informations or features. In [41] the nodes,  
 638 are geometrical region of the detector as defined by the HealPix [47]. The features of the nodes are  
 639 aggregated informations from the PMTs it contains. The edges contains geographic informations of  
 640 the nodes relative positions.

641 This data representation has the advantages to keep the topology of the detector intact. It also permit  
 642 the use of rotational invariant algorithms for the NN, thus taking advantage of the symmetries of the  
 643 detector.

644 The neural network then process the graph using Chebyshev Convolutions [48]. The performances  
 645 of the GNN are presented in figure 2.21.

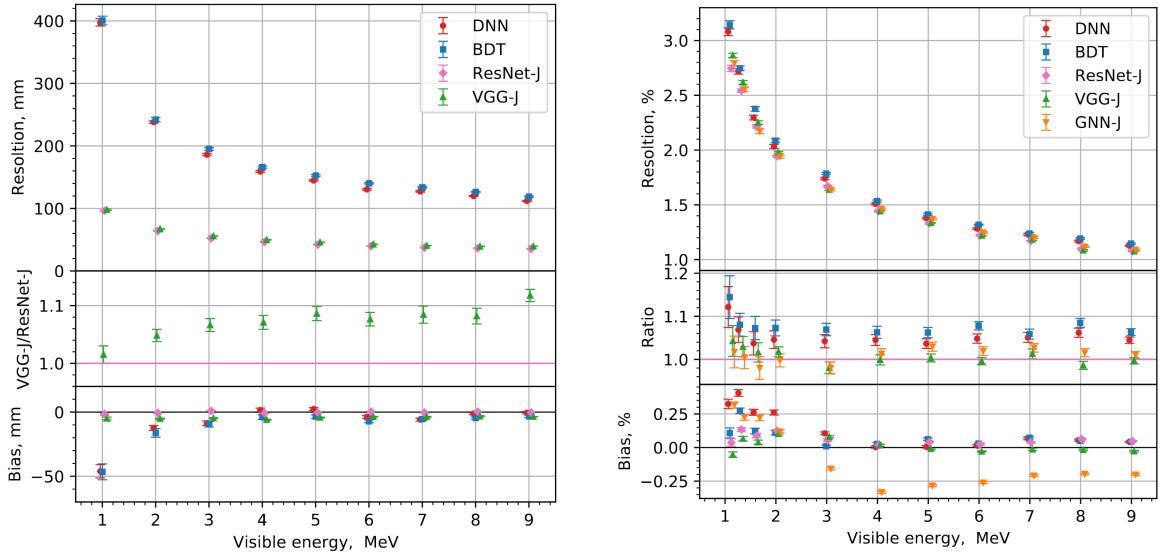


FIGURE 2.21 – Radial (left) and energy (right) resolutions of different ML algorithms. The results presented here are from [41]. DNN is a deep neural network, BDT is a BDT, ResNet-J and VGG-J are CNN and GNN-J is a GNN.

646 Overall ML algorithms show similar performances as classical algorithms in term of energy recon-  
 647 structions with the more complex structure CNN and GNN showing better performances than BDT

and DNN. For vertex reconstruction, the BDT and DNN show poor performance while CNN are on the level of the classical algorithms.

## 2.7 JUNO sensitivity to NMO and precise measurements

Now that the event have been reconstructed, selected and that the non-IBD background have been rejected, we have access to the measured energy flux from JUNO. We consider two spectra, the one measured by the LPMT system and the one measured by the SPMT system. This give rise to three possible analysis: A LPMT only analysis, a SPMT only analysis and a joint analysis. This joint analysis is the subject of the chapter 7 of this thesis.

The following details about JUNO measurement is common to the three analysis. The details and specific of the joint analysis are detailed in chapter 7.

### 2.7.1 Theoretical spectrum

To extract the oscillation parameters and the NMO from the measured spectrum, it is compared to a theoretical spectrum. This theoretical spectrum is produced based on the theory of the three flavour oscillation (see section 1.3), the measurements of the calibration and satellite experiments and Monte Carlo simulation:

- The absolute flux and the fission product fraction calibrated by TAO.
- The estimation of the neutrinos flux from other sources, such as the geoneutrinos, by theoretical model.
- The computed cross-section of  $\bar{\nu}_e$  and the LS.
- The estimation of mislabelled event, such as fast neutron events from cosmic muons, using Monte Carlo simulation.
- The measured bias and resolution of the LPMT and SPMT system by the calibration.
- The time dependent reactor parameters (age of fuel, instantaneous power of the reactors, etc...)

These systematics parameters come with their uncertainties that need to be taken into account by the fitting framework. This theoretical spectrum will, in the end, depend of the oscillation parameters of interest  $\theta_{13}$ ,  $\theta_{12}$ ,  $\Delta m_{21}^2$ ,  $\Delta m_{31}^2$ . Noise parameters can be included in the parameters spectrum such as the earth density  $\rho$  between the power plants and JUNO.

### 2.7.2 Fitting procedure

The theoretical and measured spectra are represented as two histograms depending on the energy. The theoretical spectrum is adjusted with the data using a  $\chi^2$  minimization where  $\chi^2$  is naively defined as

$$\chi^2 = \sum_i \frac{(N_{th}^i - N_{data}^i)^2}{\sigma_i^2} \quad (2.26)$$

where  $N_{th}^i$  is the number event in the  $i$ th bin of the theoretical spectrum,  $N_{data}^i$  is the number of event in the  $i$ th bin of the measured spectrum and  $\sigma_i$  is the uncertainty of this bin. Two classic statistic test exist Pearson and Neyman where the difference is the estimation of  $\sigma_i$  parameters.

This  $\sigma_i$  is composed of the systematics uncertainties discussed above but also from the statistic uncertainty of the spectrum. Considering a Poisson process, the statistic uncertainty is estimated as  $\sigma_{stat}^i = \sqrt{N^i}$ . In a Pearson test,  $N^i \equiv N_{th}^i$  whereas in a Neyman test  $N^i \equiv N_{data}^i$ . Under the assumption that the content of each bin follow a Gaussian distribution (a Poisson with high enough statistic), the two test are equivalent. But studies on Monte Carlo spectrum showed that the Pearson

and Neyman statistic are biased in opposite direction. It is easily visible where, for the same data, Pearson will prefer a higher  $N_{th}^i$  to reduce the ratio  $\frac{1}{N_{th}^i}$  whereas Neyman will prefer a lower  $N_{th}^i$  to reduce the  $(N_{th}^i - N_{data}^i)$  term.

This problematic can be circumvented by summing the two test, yielding the CNP statistic test and/or by adding a term

$$\chi^2 = \sum_i \frac{(N_{th}^i - N_{data}^i)^2}{\sigma_i^2} - \ln |\mathbf{V}| \quad (2.27)$$

where  $V$  is the covariance matrix of the theoretical spectrum yielding the PearsonV and CNPV statistic test.

The  $\chi^2$  is minimized by exploring the parameter phase space via gradient descent.

### 2.7.3 Physics results

The oscillation parameters are directly extracted from the minimization procedure and the error can be estimated directly from the procedure. For the NMO, the data are fitted under the two assumption of NO and IO. The difference in  $\chi^2$  give us the preferred ordering and the significance of our test. Latest studies show that the precision on oscillation parameters after six year of data taking will be of 0.2%, 0.3%, 0.5% and 12.1% for  $\Delta m_{31}^2$ ,  $\Delta m_{21}^2$ ,  $\sin^2 \theta_{12}$  and  $\sin^2 \theta_{13}$  respectively [11]. The expected sensitivity to mass ordering is  $3\sigma$  after 6 years [49].

## 2.8 Summary

JUNO is one the biggest new generation neutrino experiment. Its goal, the measurements of oscillation parameters with unprecedented precision and an NMO preference at the 3 sigma confidence level, needs an in depth knowledge and understanding of the detector and the physics at hand. The characterisation and calibration of the detector are of the utmost importance and the understanding of the detector response in its resolution and bias is capital to be able to correctly carry the high precision physics analysis of the neutrino oscillation.

In this thesis, I explore the usage of data-driven reconstruction methods to validate and optimize the reconstruction of IBD events in JUNO in the chapters 4, 5 and 6 and the usage of the dual calorimetry in the detection of possible mis-modelisation in the theoretical spectrum 7.



<sup>712</sup> **Chapter 3**

<sup>713</sup> **Machine learning and Artificial  
Neural Network**

<sup>715</sup> *"I have the shape of a human being and organs equivalent to those of a human being. My organs, in fact, are identical to some of those in a prostheticized human being. I have contributed artistically, literally, and scientifically to human culture as much as any human being now alive. What more can one ask?"*

Isaac Asimov, *The Complete Robot*

<sup>716</sup> Machine Learning (ML) and more specifically Neural Network (NN) are families of data-driven  
<sup>717</sup> algorithm. They are used to model complex distributions from a finite dataset to extract a generalist  
<sup>718</sup> behavior. They learn, adapt their intrinsic parameters, interactively by computing its performance  
<sup>719</sup> or loss on those dataset. They take advantage of simple microscopic operation such as *if condition* or  
<sup>720</sup> non-continuous but differentiable function like *ReLU*. Through optimizers and the combination of a  
<sup>721</sup> lot of those microscopic operations, they can obtain complex and precise behaviours.

<sup>722</sup> They are now widely used in a wide variety of domain including natural language processing,  
<sup>723</sup> computer vision, speech recognition and, the subject of this thesis, scientific studies.

<sup>724</sup> We found them in particle physics, either as the main algorithm or as secondary algorithm, for event  
<sup>725</sup> reconstruction, event classification, waveform reconstruction, etc..., domains where the underlying  
<sup>726</sup> physic and detector process is complex and highly dimensional. Physicists have traditionally been  
<sup>727</sup> forced to use simplifications or assumptions to ease the development of algorithms or equations  
<sup>728</sup> (a good example is the algorithm presented in section 2.6) where machine learning could refine and  
<sup>729</sup> take into account those effects, provided that they have enough data and computing power.

<sup>730</sup> This chapter present an overview of the different kind of machine learning methods and neural  
<sup>731</sup> networks that will be discussed in this thesis.

<sup>732</sup> **3.1 Boosted Decision Tree (BDT)**

<sup>733</sup> One of the most classic machine learning algorithm used in particle physics is Boosted Decision Tree  
<sup>734</sup> (BDT) [50] (or more recently Gradient Boosting Machine [51]). The principle of a BDT is fairly simple  
<sup>735</sup> : based on a set of observables, a serie of decisions, represented as node in a tree, are taken by the  
<sup>736</sup> algorithm. Each decision point, or node, takes its decision based on a set of trainable parameters  
<sup>737</sup> leading to a subtree of decision. The process is repeated until it reach the final node, yielding the  
<sup>738</sup> prediction. A simplistic example is given in figure 3.1.

<sup>739</sup> The training procedure follow a simple score reward procedure. During the training phase the  
<sup>740</sup> prediction of the BDT is compared to a known truth about the data. The score is then used to

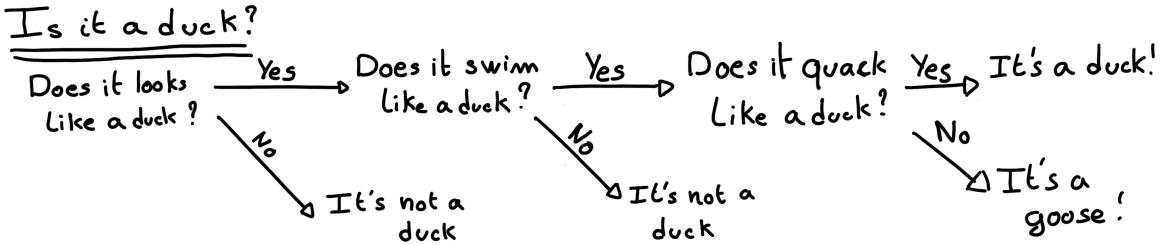


FIGURE 3.1 – Example of a BDT that determine if the given object is a duck

741 backpropagate corrections to the parameters of the tree. Modern BDT use gradient boosting where  
742 the gradient of the loss is calculated for each of the BDT parameters. Following the gradient descent,  
743 we can reach the, hopefully, global minima of the loss for our set of parameters.

## 744 3.2 Artificial Neural Network (NN)

745 One other big family of machine learning algorithm is the artificial Neural Networks (NN). The idea  
746 of developing automates which component mimic, in a simplistic way, the behavior of biological  
747 neurons emerge in 1959 with the paper “*What the Frog’s Eye Tells the Frog’s Brain*” [52]. They develop  
748 an automate where each component possess an *activation function*. Each one of those component then  
749 transmit its information to the other following a certain efficiency or *weight*. Those works influenced  
750 scientist and notably Frank Rosenblatt who published in 1958 what is considered the first neural  
751 network model the Perceptron [53].

752 Modern neural network still nowadays use the neuron metaphor to represent neural network, but  
753 approach them as a graph where the nodes are neurons possessing an activation function and edges  
754 holding the weights between those nodes. Most of the modern neural network work with the  
755 principle of neurons layers. Each neurons belong to a layer and takes input from the preceding  
756 layer and forward it result to next layer. For example the most basic set layer is the fully connected  
757 layer where each of its neurons is connected to every other neurons of the precessing layer. All  
758 the neurons posses the same activation function  $F$ . The connection between two the two layers is  
759 expressed as a tensor  $T_j^i$  where  $i$  is the index of the precedent layer and  $j$  the index of the current  
760 layer. The propagation from the layer  $I$  to  $J$  is then described as

$$J_j = F_j(T_j^i I_i + B_j) \quad (3.1)$$

761 where the learning parameters are the tensor  $T_j^i$  and the bias tensor  $B_j$ . This is the fundamental  
762 component of the Fully Connected Deep NN (FCDNN) family presented in section 3.2.1. Most of the  
763 modern neural networks use gradient descent to optimize their parameters, i.e. the gradient of the  
764 parameter  $\theta$  in respect of the loss function  $\mathcal{L}$  is subtracted to it

$$\theta_{i+1} = \theta_i - \frac{\partial \theta}{\partial \mathcal{L}} \quad (3.2)$$

765  $i$  being the training iteration index. This needs the expression of  $\mathcal{L}$  dependent of  $\theta$  to be differentiable,  
766 thus the layer and their activation function also need to be differentiable. This simple gradient  
767 descent, designated as Stochastic Gradient Descent (SGD), can be completed with first and second  
768 order momentum like with the Adam optimizer [54].

769 This description of neural networks as layer introduced the principle of *depth* and *width*, the number  
770 of layers in the NN and the number of neurons in each layer respectively.

### 771 3.2.1 Fully Connected Deep Neural Network (FCDNN)

772 Fully Connected Deep Neural Network (FCDNN) architecture is the natural evolution of the Perceptron.  
 773 The input data is represented as a first order tensor  $I_j$  and then fed forward to multiple fully  
 774 connected layers (Eq 3.1) as presented in the figure 3.2a. Most of the time, the classic ReLU function

$$\text{ReLU}(x) = \begin{cases} x & \text{if } x \geq 0 \\ 0 & \text{otherwise} \end{cases} \quad (3.3)$$

775 is used as activation function. PreLu and Sigmoid are also popular choices:

$$\text{Sigmoid}(x) = \frac{1}{1 + e^{-x}} \quad (3.4) \quad \text{PReLU}(x) = \begin{cases} x & \text{if } x \geq 0 \\ \alpha x & \text{otherwise} \end{cases} \quad (3.5)$$

776 The reasoning behind ReLU and PReLU is that with enough of them, you can mimic any continuous  
 777 function as illustrated in figure 3.2b. Sigmoid is more used in case of classification, its behavior going  
 778 hand in hand with the Cross Entropy loss function used in classification problems.

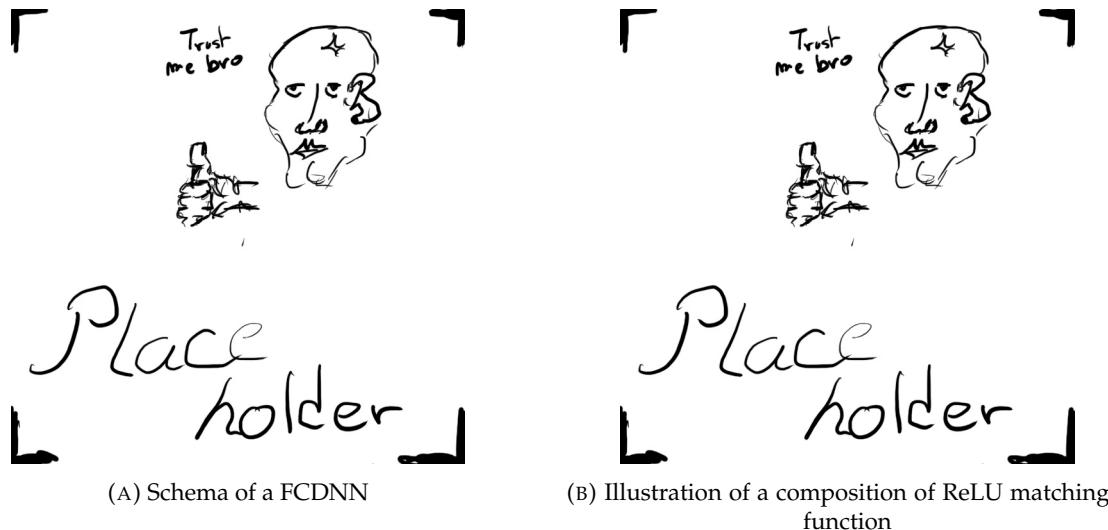


FIGURE 3.2

780 Due to its simplicity, FCDNN are also used as basic pieces for more complex architectures such as  
 781 the CNN and GNN that will be presented in the next section.

### 782 3.2.2 Convolutional Neural Network (CNN)

783 Convolutional Neural Networks are a family of neural networks that use discrete convolution filters  
 784 to process the input data. They have the advantage to be translation invariant by construction,  
 785 this mean that they are capable of detecting oriented features independently of their location on  
 786 the image. The learning parameters are located in the filters, the network thus learn the optimal  
 787 filters extract the desired feature. 2D CNN, where the filters are second order tensors that span  
 788 over third order tensors, are commonly used in image recognition [55] for classification or regression  
 789 problematics.

790 The convolution layers are commonly chained [56], reducing the input dimension while increasing  
 791 the number of filters. The idea behind is that the first layers will process local informations and the  
 792 latest layers will process more global informations. To try to preserve the amount of information, we

793 tend double the numbers of filters for each division of the input data. The results of the convolution  
 794 filters is commonly then flattened and feed to a smaller FCDNN which will process the filters results  
 795 to yield the desired output.

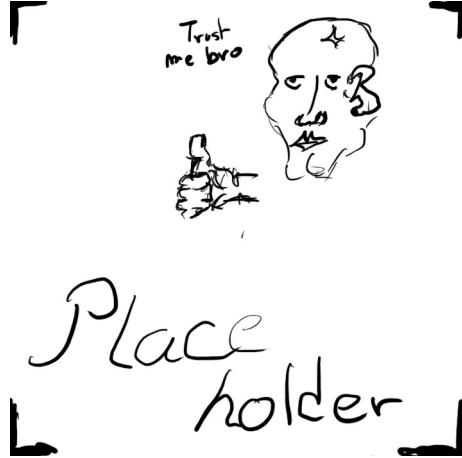


FIGURE 3.3 – Illustration of a CNN convolution filter

### 796 3.2.3 Graph Neural Network (GNN)

797 Graph neural network is a family of neural network where the data is represented as a graph  $G(\mathcal{N}, \mathcal{E})$   
 798 composed of vertex or node  $n \in \mathcal{N}$  and edges  $e \in \mathcal{E}$ . The edges are associated to two nodes  $(u, v) \in$   
 799  $\mathcal{N}^2$ , “connecting” them. The node and the edges can hold features, commonly represented as vector  
 800  $n \in \mathbb{R}^{k_n}$ ,  $e \in \mathbb{R}^{k_e}$ . We can thus define a graph using two tensors  $A_e^{ij}$  the adjacency tensors that hold  
 801 the features  $e$  of the edge connecting the node  $i$  and  $j$  and the tensor  $N_v^i$  that hold the features  $v$  of a  
 802 node  $i$ .

803 To efficiently manipulate such object we need to structurally encode their property in the neural  
 804 network architecture: each node is equivalent (as opposite to ordered data in a vector), each node has  
 805 a set of neighbours, ... One of this method is the message passing algorithm presented historically  
 806 in “Neural Message Passing for Quantum Chemistry” [57]. In this algorithm, with each layer of  
 807 message passing a new set of features is computed for each node following

$$n_i^{k+1} = \phi_u(n_i^k, \square_j \phi_m(n_i^k, n_j^k, e_{ij}^k)); n_j \in \mathcal{N}'_i \quad (3.6)$$

808 where  $\phi_u$  is a differentiable update function,  $\square_j$  is a differentiable aggregation function and  $\phi_m$  is a  
 809 differentiable message function.  $\mathcal{N}'_i = \{n_j \in \mathcal{N} | (n_i, n_j) \in \mathcal{E}\}$  is the set of neighbours of  $n_i$ , i.e. the  
 810 nodes  $n_j$  from which it exist an edge  $e_{i,j} \rightarrow (n_i, n_j)$ .  $k$  is the layer on which the message passing  
 811 algorithm is applied.  $\square$  need also a few other property if we want to keep the graph property, most  
 812 notably the permutational invariance of its parameters (example: mean, std, sum, ...).

813 The edges features can also be updated, either by directly taking the results of  $\phi_m$  or by using another  
 814 message function  $\phi_e$ .

815 Message passing is a very generic way of describing the process of GNN and it can be specialized  
 816 for convolutional filtering [48], diffusion [58] and many other specific operation. GNN are used in a  
 817 wide variety of application such as regression problematics, node classification, edge classification,  
 818 node and edge prediction, ...

819 It is a very versatile but complex tool.

**3.2.4 Adversarial Neural Network (ANN)**

The adversarial machine learning, Adversarial Neural Networks (ANN) in the case of neural network, is a family of unsupervised machine learning algorithms where the learning algorithm (generator) is competing against another algorithm (discriminator). Taking the example of Generative Adversarial Networks, concept initially developed by Goodfellow et al. [59], the discriminator goal is to discriminate between data coming from a reference dataset and data produced by the generator. The generator goal, on the other hand, is to produce data that the discriminator would not be able to differentiate from data from the reference dataset. The expression of duality between the two models is represented in the loss where, at least a part of it, is driven by the results of the discriminator.



<sup>829</sup> **Chapter 4**

<sup>830</sup> **Image recognition for IBD  
reconstruction with the SPMT system**

<sup>832</sup> As explained in chapter 2, JUNO is an experiment composed of two systems, the Large Photo Tube  
<sup>833</sup> Multiplier (LPMT) and the Small Photo Multiplier Tubes (SPMT). Both of the system observe the  
<sup>834</sup> same physics event inside of the same medium but they differ in their photo-coverage, respectively  
<sup>835</sup> 75.2% and 2.7%, their dynamic range, a thousands versus a few dozen, and their back-end electronics  
<sup>836</sup> (see section 2.2.2).

<sup>837</sup> Due to their differences they are complementary and their strength and weakness. One important  
<sup>838</sup> point the difference in expected resolution, the LPMT system outperform largely the SPMT system  
<sup>839</sup> but is subject to effects such as saturation [28] that could bias the reconstruction, effect that the SPMT  
<sup>840</sup> system is impervious to. Also, due to the dynamic range of the LPMT, in case of high energy and  
<sup>841</sup> high density event such as core-collapse supernova, the LPMT system could saturate and the lack of  
<sup>842</sup> photo-coverage become a benefit.

<sup>843</sup> Thus it is important to have reconstruction and analysis combined but also dedicated to each sys-  
<sup>844</sup> tem, taking into account the specificity. The subject of this work is to propose a machine learning  
<sup>845</sup> algorithm for the SPMT reconstruction based on Convolutional Neural Network (CNN).

<sup>846</sup> **4.1 Motivations**

- <sup>847</sup> — Promise of machine learning -> Exploit raw data
- <sup>848</sup> — Victor already done reco for SPMT
- <sup>849</sup> — Can CNN give similar results ? Better results ?
- <sup>850</sup> — Multiple reco methods good for reconstruction
- <sup>851</sup> — Comparison, difference in behavior

<sup>852</sup> As explained in chapter 3, Machine Learning (ML) algorithms shine when modeling complex distri-  
<sup>853</sup> bution from a given dataset. In our case, we have access to complete monte-carlo simulation of our  
<sup>854</sup> detector to produce arbitrary large datasets that could represent multiple years of data taking. Also,  
<sup>855</sup> due to the nature of monte-carlo simulation we have access to exact truth of the modelled event like  
<sup>856</sup> the position, energy and nature of the particle.

<sup>857</sup> One of expectation when using data-driven, and so, machine learning algorithm is that the algorithm  
<sup>858</sup> will be able to use the entirety of the raw information, preventing any loss in precision due to miss-  
<sup>859</sup> modeling or simplification of the underlying physic process.

<sup>860</sup> Another algorithm was developed in the collaboration for energy and vertex reconstruction [60] and  
<sup>861</sup> will serve as a performance reference in this work. The two methods will also be studied event by  
<sup>862</sup> event to try to understand and unveil coherence or incoherence between their performances.

## 4.2 Method and model

- JUNO is an array of sensor following a quasi uniform and isotropic geometric repartition -> Basically pixels -> Image
- CNN is good for image processing (cite a lot of things)
- Details the architecture (Inspired from VGG 16)
  - Convolutional layers
  - Pooling layers -> Twice the channels when pooling by 2 -> Keep the same "amount" of information
  - Dropout (introduce overtraining, maybe introduce overfitting in ML chapter ?)
  - Vectorization fed to FCDNN

One of simplest way to look at JUNO data is to consider the detector as an array of geometrically distributed sensors on a sphere. Their repartition is almost homogeneous, on this sphere surface providing an almost equal amount of information per unit surface on this sphere. One kind of data representation that is ordered, presenting the same property as a plane is the image, which is a bounded discrete euclidian plane.

Using this representation we can arrange the SPMTs on this plane while keeping most of the neighbouring informations. The pros and cons of this representation is detailed in section 4.2.2.

The most common approach in machine learning for image processing and image recognition is the Convolutional Neural Network (CNN). It is widely used in research and industry [56, 61–63] due to its strengths (see section 3.2.2) and has proven its relevance in image processing.

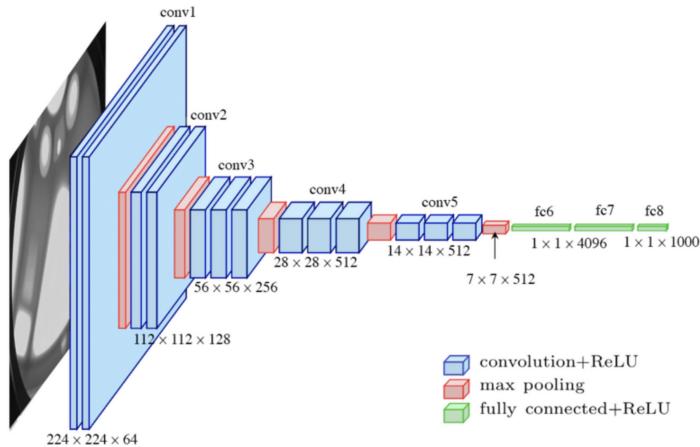


FIGURE 4.1 – Graphic representation of the VGG-16 architecture, presenting the different kind of layer composing the architecture.

### 4.2.1 Model

The architecture we use is derived from the VGG-16 architecture [56] illustrated in figure 4.1. We define a set of hyperparameters which values are presented in table 4.1.  $N_{blocks}$  is the number of convolution blocks, a block being composed of two convolutional layers with  $3 \times 3$  filters using ReLU activation function, a  $3 \times 3$  max-pooling layer (except for the last block) and a dropout layer. The number of channels is dependent of the depth of the block  $N_{channels}^i = i * N_{channels}$ ,  $i \in [1..N_{blocks}]$ .

The result of the last convolution layer is then flattened then fed forward to the FCDNN. Its configuration is expressed as a sequential of fully connected layer using the PReLU activation function. For example  $2 * 1024 + 2 * 512$  is the sequence of 2 layers with a width of 1024 followed by 2 other layers

892 with a width of 512. Finally the last layer is 4 neurons wide without activation function, one for each  
 893 component of the interaction vertex: Energy, X, Y, Z.

We explore in this work two different activation functions

$$(E + V)(E, x, y, z) = \left\langle (E - E_{true})^2 + 0.85 \sum_{\lambda \in [x, y, z]} (\lambda - \lambda_{true})^2 \right\rangle \quad (4.1)$$

$$(E_r + V_r)(E, x, y, z) = \left\langle \frac{(E - E_{true})^2}{E_{true}} + \frac{10}{R} \sum_{\lambda \in [x, y, z]} (\lambda - \lambda_{true})^2 \right\rangle \quad (4.2)$$

894 where  $R$  is the radius of the CD. With the energy in MeV and the distance in meters, we use the factor  
 895 0.85 and 10 to equilibrate the two term of the loss function so they have the same magnitude.

896 — The loss function  $(E + V)$  is close to a simple Mean Squared Error (MSE). MSE is one of the  
 897 most basic loss function, the derivative is simple and continuous in every point. It is a strong  
 898 starting point to explore the possibility of CNNs.

899 —  $(E_r + V_r)$  can be seen as a relative MSE.

900 The idea is that: due to the inherent statistic uncertainty over the number of collected Number of  
 901 Photo Electrons (NPE), the absolute resolution  $\sigma(E - E_{true})$  will be larger at higher energy than at  
 902 low energy. But we expect the *relative* energy resolution  $\frac{\sigma(E - E_{true})}{E_{true}}$  to be smaller at high energy than  
 903 lower energy as illustrated in figure 2.19. Because of this, by using simple MSE the most important  
 904 part in the loss come from the high energy part of the dataset whereas with a relative MSE, the most  
 905 important become the low energy events in the dataset. We hope that by using a relative MSE, the  
 906 neural network will focus on low energy events where the reconstruction is considered the hardest  
 907 part of the dataset.

$N_{blocks}$	2, 3, 4
$N_{channels}$	32, 64, 128
FCDNN configuration	2 * 1024 2 * 2048 + 2 * 1024 3 * 2048 + 3 * 512 2 * 4096
Loss	$E + V, E_r + V_r$

TABLE 4.1 – Hyperparameters used for the CNN

908 Each one of the combination of those hyperparameters is then tested and compared to each other  
 909 over an analysis sample. We cannot use the mean loss because we consider multiple loss function so  
 910 we use multiple observables:

- 911 — The mean absolute energy error  $\langle E \rangle = \langle |E - E_{true}| \rangle$ . It is an indicator of the energy bias of our  
 912 reconstruction.
- 913 — The standard deviation of the energy error  $\sigma E = \sigma(E - E_{true})$ . This the indicator on our  
 914 precision in energy reconstruction.
- 915 — The mean distance between the reconstructed vertex and the true vertex  $\langle V \rangle = \langle |\vec{V} - \vec{V}_{true}| \rangle$ .  
 916 This an indicator of the bias and precision of our vertex reconstruction.
- 917 — The standard deviation of the distance between the true and reconstructed vertex  $\sigma V = \sigma |\vec{V} - \vec{V}_{true}|$ .  
 918 This is an indicator if the precision in our vertex reconstruction.

919 For the vertex indicator, it is worth noting that  $\langle V \rangle$  is not expected to be zero and so if unfit to be  
 920 used for the loss. Indeed, let's say we consider error on each of the component as random variable  
 921 following a normal distribution. We allow ourself to use this representation as our signal possess a  
 922 strong statistical uncertainty in NPE that follow a Poisson law, i.e. a Gaussian law  $\mathcal{N}$  when NPE is

923 high enough which is the case for our signal. So following:

$$\Delta V = |\vec{V} - \vec{V}_{true}| = \sqrt{\Delta X^2 + \Delta Y^2 + \Delta Z^2}; \Delta X, \Delta Y, \Delta Z \sim \mathcal{N} \quad (4.3)$$

924 then

$$\Delta V \sim \chi \quad (4.4)$$

925 where  $\chi$  is a Chi law which probability density function is different from 0 only in  $\mathbb{R}^+$

## 926 4.2.2 Data representation

- 927 — Data is 240x240 images
- 928 — Following  $\theta$  and  $\phi$  distribution, explain the coordinate system of JUNO
- 929 — Optimized for  $\approx 1$  SPMT/pixel
- 930 — 1 Charge channel
- 931 — 1 Time channel
- 932 — Discuss data format
- 933 — Empty pixel ?  $\rightarrow Q = 0, T = 0$ , what does it means/says ? 0 = no signal in a way
- 934 — Image distortion
  - 935 — **Maybe speak of this in the conclusion ?** Could be done in two step:
  - 936 — 1. Reconstruct  $\theta$  and  $\phi$
  - 937 — 2. "Rotate" the image so the event is at the center of the image  $\rightarrow$  Prevent distortion +
  - 938 — reconstruction E and R become pseudo rotational invariant (as they should be)
- 939 — 1 Millions MC e+ events for training (900k for train, 50k for validation and 50k for test)
- 940 — MC for the moment, will need to retrain with mix of calibration data (Good question,
- 941 — is the CNN PID agnostic ?)
- 942 — 47 IBD/day  $\rightarrow$  1M event is 21k days of data (for reference, 6 years of data is 94k events)
- 943 — Events are "optimistic"
  - 944 — No pile-up
  - 945 — w/o neutrons
  - 946 — time window is decided by electronics
  - 947 — We want to reconstruct the E from  $\bar{\nu}_e$
  - 948 — Difference between multiple E  $\rightarrow E_{vis}, E_{rec}, E_k$

949 The data used during this analysis is monte carlo data using the official JUNO simulation software  
 950 (see section 2.5 for details). The simulated data is composed of positron events, uniformly distributed  
 951 in the CD volume and in kinetic energy over  $E_k \in [0; 9]$  MeV producing a deposited energy  $E_{dep} \in$   
 952  $[1.022; 10.022]$  MeV. This is done to mimic the signal produced by the IBD prompt signal. Uniform  
 953 distribution are used so that the CNN does not learn a potential energy distribution, favoring some  
 954 part of the energy spectrum instead of other.

955 This data is represented as  $240 \times 240$  images, equivalent to third order tensor, with a charge  $Q$  channel  
 956 and a time  $t$  channel. The SPMTs are then projected on the plane following illustrated in figure 4.2a,  
 957 the  $x$  position is proportional to  $\theta$  and the  $y$  position is defined by  $\phi \sin \theta$  in spherical coordinates.  
 958  $\theta = 0$  is defined as being the top of the detector and  $\phi = 0$  is defined as an arbitrary direction in  
 959 the detector. In practice, this is the  $\phi = 0$  given by the MC simulation.

960 When two SPMTs are in the same pixel, the charges are summed and the lowest of the hit-time  
 961 is chosen. This choice come from the fact that we expect the photon propagation to be uniform,  
 962 there should no physical meaning in two signal photons coming one after the other except that the  
 963 two PMT are at different position. The SPMTs being located close to each other, the reduction of  
 964 information into the first hit time should not impact to much the reconstruction, at least less than the  
 965 intrinsic time resolution and statistic uncertainty. The only potential problem in using this first time  
 966 come from the Dark Noise (DN). Its time distribution is uniform over the signal and could come

before a signal hit on the other SPMT in the pixel. In that case, the time information in the pixel become irrelevant and we lose the time information for this part of the detector. As illustrated in figure 4.2a the dimension have been chosen optimized so that at most two SPMTs are in the same pixel while keeping the number of empty pixels relatively low to prevent this kind of issue.

While it could be possible to use bigger images to prevent overlapping, keeping image small images gives multiple advantages:

- As presented in section 4.2.1, the convolution filter we use are  $3 \times 3$  convolution filter, meaning that if SPMTs would be separated by more than one pixel, the first filter would only see one SPMT per filter. This behavior would be kind of counterproductive as the first convolution block would basically be a transmission layer and would just induce noise in the data.
- It keep the network relatively small, while this do not impact the convolution layers, the flatten operation just before the FCDNN make the number parameters in the first layer of it dependent on the size of the image.
- It reduce the number of empty pixel in the image.

The question of empty pixel is an important question in this data representation. Their is two kind of empty pixel in the data.

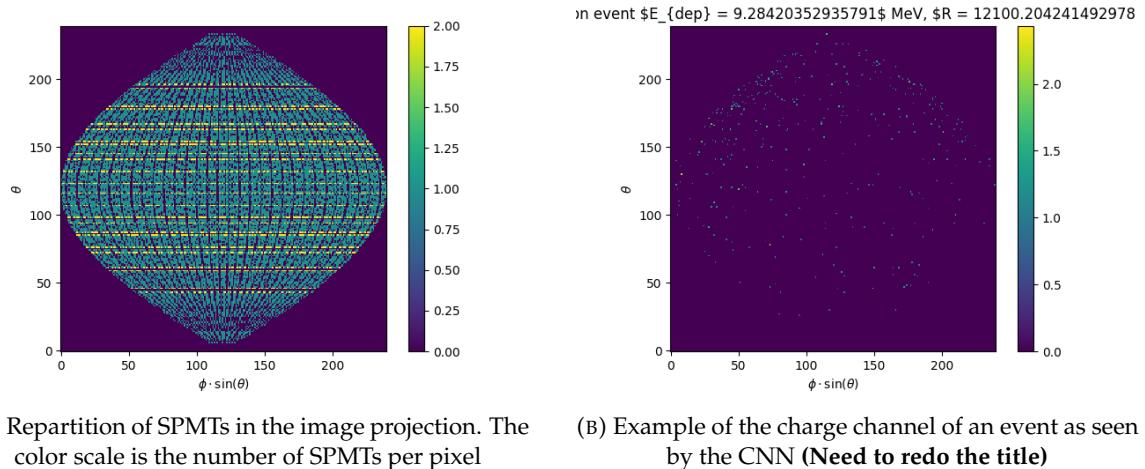


FIGURE 4.2

The first kind is pixel that contain a SPMT but the SPMT did not get hit nor registered any dark noise during the event. In this case, the charge layer is zero, which have a physical meaning but then come the question of the time layer. One could argue that the correct time would be infinity (or the largest number our memory allows us) because the hit “never” happened, so extremely far from the time of the event. This cause numerical problem as large number, in the linear operation that are happening in the convolution layers, are more signifiant than smaller value. We could try to encode this feature in another way but no number have any significance due to our time being relative to the trigger of the experiment so  $-1$  for example is out of question. Float and Double gives us access to special value such as NaN (Not a Number) [64] but the behavior is to propagate the NaN which leaves us with NaN for energy and position. We choose to keep the value 0 because it’s the absorbing element of multiplication and can be though as no activation in the ReLU activation function.

The second kind of pixel is pixel that do not represent parts of the detector such as the corners of the images. The question is basically the same, what to put in the charge and the time channel. The decision is to set the charge and time at 0 following the reasoning presented above. Its important to keep in mind that the fact that a part of the detector that has not been hit is also an information: There is no signal in this part of the detector. This problematic will be explored in more details in chapter 5.

1000 Another problematic that happens with this representation, and this is not dependent of the chosen  
1001 projection, is the deformation in the edges of the image and the loss of the neighbouring information  
1002 in the for the SPMTs at the edge of the image  $\phi \sim 180^\circ$ . This deformation and neighbouring loss  
1003 could be partially circumvented as explained in section 4.4

## 1004 4.3 Results

- 1005 — Comparison with victor results  
1006 — **More details when I'll look into the retrained data**  
1007 — Discuss of the differences  
1008 — Discuss of the principle of error decorrelation  
1009 — Possible improvements  
1010 — Combining algorithms  
1011 — Average sum

## 1012 4.4 Prospect

## 1013 4.5 Conclusion

1014 Intoduction next chapter

1015 **Chapter 5**

1016 **Graph representation of JUNO for IBD  
reconstruction with the LPMT system**

1017



<sup>1018</sup> **Chapter 6**<sup>1019</sup> **Reliability of machine learning  
methods**  
<sup>1020</sup>

1021

*"Psychohistory was the quintessence of sociology; it was the science of human behavior reduced to mathematical equations. The individual human being is unpredictable, but the reactions of human mobs, Seldon found, could be treated statistically"*

*Isaac Asimov, Second Foundation*



<sup>1022</sup> **Chapter 7**

<sup>1023</sup> **Joint fit between the SPMT and LPMT  
spectra**

<sup>1024</sup>



<sup>1025</sup> **Chapter 8**

<sup>1026</sup> **Conclusion**



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