

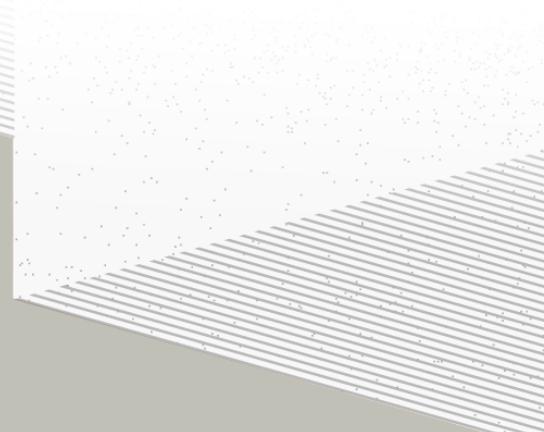
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Par

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**Deep learning methods and Dual Calorimetric analysis for high precision neutrino oscillation measurements at JUNO**

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<sup>3</sup> A mon Père  $\hat{\text{D}}$

<sup>4</sup> A ma Mère  $\hat{\text{A}}$



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<sup>125</sup> **Remerciements**



# <sup>126</sup> Introduction

<sup>127</sup> The Standard Model of particle physics (SM) has been remarkably successful at accounting for,  
<sup>128</sup> or predicting experimental observations in the laboratory. However, it is the subject of several  
<sup>129</sup> limitations. For instance, it provides a mechanism to explain the existence of mass but can't predict  
<sup>130</sup> the peculiar pattern followed by fermion masses. The same applies to CP violation. The SM predicts  
<sup>131</sup> its existence but not the amplitude necessary to explain the baryonic asymmetry of the Universe. For  
<sup>132</sup> such reasons, one can assume the SM is the manifestation of a more fundamental physics, Beyond  
<sup>133</sup> the Standard Model (BSM).

<sup>134</sup> Neutrino physics is a window on BSM. Indeed, the mass of known neutrinos is at least 5 order of  
<sup>135</sup> magnitudes below that of the lightest fermion, which further deepens the issue of fermion mass  
<sup>136</sup> generation. Some solutions have implication on the nature of neutrinos – dirac or majorana fermions  
<sup>137</sup> ? – which one of the big unknowns in this domain. Additional neutrinos beyond the three presently  
<sup>138</sup> known shall also be considered. The way neutrinos mix flavor to make neutrino oscillation possible  
<sup>139</sup> is also unexplained. This is one of the tasks of BSM models to answer such questions. Before that, a  
<sup>140</sup> good part of the World experimental program in the 10 coming years is to complete the exploration  
<sup>141</sup> of 3-neutrino physics by answering mainly two questions : does CP violation exist in the lepton  
<sup>142</sup> sector ? What is the Neutrino Mass ordering (NMO) ? An introduction to neutrino physics will be  
<sup>143</sup> given in Chapter 1.

<sup>144</sup>

<sup>145</sup> The Jiangmen Underground Neutrino Observatory (JUNO), currently under construction in China,  
<sup>146</sup> aims to address these questions, particularly the determination of the NMO. JUNO's approach is  
<sup>147</sup> to study reactor antineutrinos emitted from nearby nuclear power plants. By precisely measuring  
<sup>148</sup> the energy spectrum of these antineutrinos after oscillation, JUNO seeks to detect the subtle inter-  
<sup>149</sup> ference patterns in the spectrum that are sensitive to the NMO. The ability to achieve this requires  
<sup>150</sup> unprecedented precision in both the energy resolution and the calibration of the detector's response  
<sup>151</sup> to neutrino events. JUNO is expected to start data collection in 2025, with the goal of determining the  
<sup>152</sup> NMO at a significance level of  $3-4\sigma$  after six years of data taking. At the heart of JUNO's experimental  
<sup>153</sup> design is its dual calorimetry system, comprising two separate sets of photomultipliers-large (LPMT)  
<sup>154</sup> and small (SPMT) PMTs that allow for independent energy measurements of the same events. This  
<sup>155</sup> dual system is not only essential for improving energy resolution but also for providing cross-checks  
<sup>156</sup> that ensure systematic uncertainties are well-understood and minimized. Achieving JUNO's goals  
<sup>157</sup> depends on this dual calorimetry system, as it will enable precise reconstruction of the energy  
<sup>158</sup> spectrum and the identification of potential discrepancies between the two systems.

<sup>159</sup>

<sup>160</sup> Another emerging area of importance in particle physics experiments is the application of machine  
<sup>161</sup> learning (ML) techniques. Over the past decade, ML methods, particularly deep learning, have been  
<sup>162</sup> increasingly used to tackle complex problems in event classification, reconstruction, and even data  
<sup>163</sup> generation like the High luminosity LHC Upgraded experiments. Performant online reconstruction,  
<sup>164</sup> critical for the trigger systems of such experiments, is another example. The complexity of the data  
<sup>165</sup> and the required precision in experiments such as JUNO make ML an attractive tool. In particular,  
<sup>166</sup> Neural Networks (NNs) and other advanced ML models have shown potential for improving the  
<sup>167</sup> accuracy of energy reconstruction and other key analysis tasks. However, for the results obtained

168 using ML methods to be trusted by the scientific community, the reliability of these methods must be  
169 rigorously demonstrated. An introduction to ML, and in particular Neural Network (NN) is given  
170 in Chapter 3.

171  
172 This thesis was performed in the framework of the Neutrino group at Subatech, since October  
173 2021. The exploratory works reported in this manuscript addresses the subjects mentioned above,  
174 in the particular context of the measurement by JUNO of the reactor antineutrino oscillation to  
175 determine the NMO. Before the start of this thesis, several ML energy reconstruction algorithms  
176 – Boosted Decision Trees (BDT), Fully Connected Neural Networks (FCNN), Convolutional Neu-  
177 ral Networks (CNNs) and Graph Neural Networks (GNNs) – had already been developed within  
178 the collaboration. Their performance seems to match that of the classical algorithm but not to do  
179 convincingly better. We have explored a possibility to do better by developing a GNN with an  
180 innovative architecture tailored to the JUNO experiment. Before that, we developed a CNN for the  
181 reconstruction of the anti-neutrino energy using only JUNO’s small PMTs system. This CNN is  
182 useful in particular in Chapter 7 as there is official SPMT only reconstruction in the collaboration yet.  
183 These algorithms are described in Chapters 4 and 5.

184 We have been the first in JUNO to address the issue of ML reliability. We have followed two paths  
185 for that. First, a simple approach is to compare event per event the results obtained by various algo-  
186 rithms, to find discrepancies, and more generally differences or common points in the way detector’s  
187 information is used. This requires to implement in JUNO’s official software algorithms traditionally  
188 developed standalone, as well as the necessary software tools. This was our contribution there. The  
189 second path was to explore the feasibility of an Adversarial Neural Network (ANN) to generate  
190 (and therefore identify) scenarios of discrepancies between raw data in the real detector and in the  
191 detector’s simulation. The focus here is on discrepancies that could alter JUNO’s results on NMO, but  
192 are too subtle to be detected via usual data/MC comparisons in control samples. This is presented  
193 in Chapter 6.

194  
195 We have already mentioned earlier it is crucial for JUNO to understand its energy scale with a  
196 good precision. This is the raison d’être of the existence of two calorimetric readout systems : the  
197 large (LPMT) and small (SPMT) photomultipliers systems. It allows Dual Calorimetry techniques  
198 to constrain our understanding of the reconstruction. The last subject of this thesis explores for the  
199 first time one of them : the Dual Calorimetry with neutrino oscillation, which leverages potential  
200 discrepancies between the oscillation analyses performed with each system. Our work on this is  
201 described in Chapter 7. It was also the occasion of technical developments on the analysis framework  
202 used at Subatech. These improvements will be very useful for future analyses of the group, beyond  
203 Dual calorimetry.

<sup>204</sup> **Chapter 1**

<sup>205</sup> **Neutrino physics**

*I have done a terrible thing, I have postulated a particle that cannot be detected.*

Wolfgang Pauli – “Foreword” by Frederick Reines to “Spaceship Neutrino” by Christine Sutton, (p. xi), 1992.

<sup>207</sup> **Contents**

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<sup>220</sup> Our understanding of the universe describes it as composed of elementary components called elementary particles; the study of these particles is therefore particle physics. The best established <sup>221</sup> theoretical model describing these particles and their interactions is the Standard Model (SM). The <sup>222</sup> SM has successfully described many phenomena observed in particle physics over the past decades. <sup>223</sup> However a certain number of limitations affect the SM and suggest it is the manifestation at presently <sup>224</sup> accessible energies of a more fundamental physics, which we call Physics Beyond the SM (BSM). <sup>225</sup>

<sup>226</sup> In this chapter, I describe briefly the Standard Model and its limitations in Section 1.1, then delve a <sup>227</sup> bit further into the specifics of neutrino physics in Section 1.2.

<sup>228</sup> **1.1 Introduction to the Standard model**

<sup>229</sup> The SM categorizes elementary particles into two categories: the *fermions* constituting matter and <sup>230</sup> the *bosons* that mediate their interactions. The fermions are themselves divided into two categories, <sup>231</sup> the *quarks* and the *leptons*. Figure 1.1 shows the elementary particles and their classification. Each <sup>232</sup> one of these particles is characterized by the value of their quantum numbers, the main ones being <sup>233</sup> their mass  $m$ , spin  $J$ , electric charge  $Q$ , and the quantum numbers playing an analogue role for <sup>234</sup> the weak (weak isospin) and strong interactions (color). The leptons also possess a leptonic quantum <sup>235</sup> number  $L = 1$  and a flavor quantum number  $L_{e,\mu,\tau}$  corresponding to their family: electronic, muonic, <sup>236</sup> or tauonic. The leptons are thus split into three families: the electronic  $L_e = 1 \rightarrow (e, \nu_e)$ , muon <sup>237</sup>  $L_\mu = 1 \rightarrow (\mu, \nu_\mu)$ , and tau  $L_\tau = 1 \rightarrow (\tau, \nu_\tau)$  families, each composed of a charged particle  $Q = 1$  and

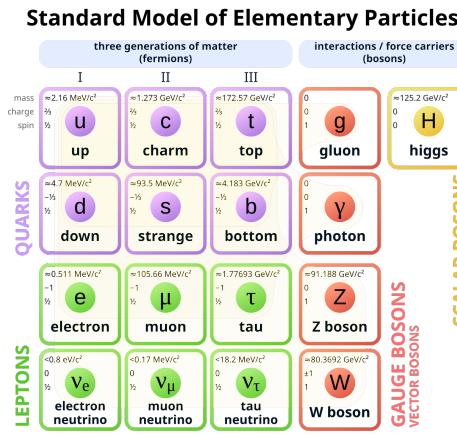


FIGURE 1.1 – List of the elementary particles in the Standard Model. The antiparticles are not displayed.

238 a neutral particle  $Q = 0$ . The neutral leptons are named the *neutrinos*, represented by the character  
 239  $\nu$ .

240 Each fermion also possesses an antiparticle of opposite charges and opposite leptonic and flavor  
 241 quantum numbers. Thus, the antiparticle of the electron  $e(Q = 1, L = 1, L_e = 1)$ , the positron, is  
 242 defined as  $e^+(Q = -1, L = -1, L_e = -1)$ .

243 The particles of the SM interact with each other via four interactions or forces. Three of these are  
 244 described by the SM through the exchange of a boson:

- The strong force, described by the exchange of a gluon. Only quarks are sensitive to it. This force is very short-range,  $\sim 10^{-15}$  m, the size of a nucleus. It's the strong force that allows the cohesion of nuclei inside atoms. As its name indicates, it is the strongest of the four interactions. This interaction has two important properties. It increase with the distance, causing quark confinement: the quarks are never observed individually but uncoloured quark combinations, the hadrons, are not affected. It decrease at short distance, creating the asymptotic freedom: the interaction strength become arbitrarily small as the distance between particles decrease.
- The electromagnetic force is described by the exchange of a photon. This force has unlimited range, and all the charged particles – quarks and charged leptons – are sensitive to it. It is responsible for every electro-magnetic effect, like the bonding of electrons to the nucleus. Its relative strength compared to the strong force is 1/137.
- The weak force, carried by the  $Z^0$  and  $W^\pm$  bosons. Every fermion is sensitive to it. Its range is  $\sim 10^{-18}$  m, about 0.1% the size of a proton. Its relative strength to the strong force is  $10^{-6}$ , explaining its name. It is responsible, for instance, of nuclear beta decay many particles decays. We distinguish two types of weak interaction: through neutral currents – exchange of a  $Z^0$  – and charged currents – exchange of a  $W^\pm$  boson.

261 The final force, not described by the Standard Model, is the gravitational force. Its range is infinite  
 262 and concerns every massive ( $m \neq 0$ ) particle. Its relative strength compared to the strong force is  
 263  $6 \times 10^{-39}$ . Extensions to the SM propose a supplementary boson, the graviton, that would be the  
 264 carrier of the gravitational force, but it has yet to be detected [1, 2].

### 265 1.1.1 Interactions and symmetries

266 Symmetries are fundamental components of modern particle physics. As described in Noether's  
 267 theorem [3], the invariance or non-invariance of the physics laws under transformations (translation,

rotation, etc.), represented by the formal invariance of the SM Lagrangian  $\mathcal{L}$  under those transformations, implies the conservation of a quantity. This Lagrangian is the mathematical object that allows us to make quantitative predictions. Following here again Noether's theorem after having observed that interactions seem to conserve quantities (like the electric charge or the weak isospin), its form is guided by the necessity to be invariant under a local gauge transformation, namely under the  $SU(3)_C * SU(2)_L * U(1)_Y$  gauge group.

The invariance of  $\mathcal{L}$  under translation in space characterizes the conservation of the momentum, while rotational invariance in space leads to the conservation of the angular momentum and invariance under translation in time the conservation of energy, etc. If the transformation is continuous, the sum of the quantum numbers is conserved in an interaction.

Invariances under discrete transformations also conserve quantum numbers. Three discrete transformations are important for the SM:

- The parity  $P$  symmetry transforms  $(\vec{x}, t) \rightarrow (-\vec{x}, t)$ , reversing the handedness of space. The momentum therefore becomes  $\vec{p} \rightarrow -\vec{p}$  and the helicity  $\frac{\vec{p} \cdot \vec{s}}{|\vec{p}|}$ , where  $\vec{s}$  is the spin, changes sign.
- Time reversal  $T$  where  $(\vec{x}, t) \rightarrow (\vec{x}, -t)$ , inverting the initial and final states of an interaction. For example,  $A + B \rightarrow C$  becomes  $C \rightarrow A + B$ . The momentum and the spin both change sign, leaving the helicity unchanged.
- Charge conjugation  $C$ , which replaces the particles by their antiparticles counterpart and vice-versa, leaving the momentum, spin and helicity unchanged.

The  $C, P$  and their combination  $CP$  symmetry were believed to be conserved until 1956, the discovery of their violation [4–6] in weak interactions revealed the non-triviality of its nature.

The strong and electromagnetic interactions are invariant under all discrete and combined  $CP$  transformations. The weak interaction is only invariant under  $CPT$ .

The fundamental symmetry  $CPT$  – the combination of  $C, P$  and  $T$  symmetries –, is an exact symmetry in the SM. This means that any process where the particles are switched with their anti-particles, their helicity is reversed and initial and final state are swapped, must occur with the same probability than the initial process. This implies that the mass, life times, absolute values of electric charge and magnetic moments of particles and antiparticles must be the same. To date, there's no experimental sign of  $CPT$  violation.

## 297 Higgs mechanism

298 The introduction of fermion and boson masses in the SM Lagrangian requires a spontaneous breaking  
 299 of the  $SU(2)_L * U(1)_Y$  gauge symmetry, which is achieved through the Higgs mechanism. This  
 300 process introduces the Higgs field, whose non-zero vacuum expectation value breaks the electroweak  
 301 symmetry, giving masses to the  $W$  and  $Z$  bosons while keeping the photon massless, thus pre-  
 302 serving electromagnetism. The introduction of the Higgs field in the Yukawa potential of the SM  
 303 Lagrangian makes apparent a fermion mass term, resulting from the coupling of the Higgs field  
 304 with the fermions. This term is proportional to a Higgs-fermion coupling constant, which value  
 305 is unpredicted by the theory. The theory, proposed by Peter Higgs, François Englert, and Robert  
 306 Brout in 1964 [7–10], was experimentally confirmed with the discovery of the Higgs boson by the  
 307 ATLAS and CMS collaborations at CERN in 2012 [11, 12].

## 308 Flavor Mixing

309 The phenomena of flavor mixing is when the flavor eigenstate differ from the mass eigenstate –  
 310 the interacting state is a superposition of mass state. In the quark sector, flavor mixing has been  
 311 observed in strangeness (one of the quark quantum numbers) violating decay, like  $K^+ \rightarrow \mu^+ + \nu_\mu$ .

312 The explanation is that the quark state that undergo the decay, i.e. the weak interaction, was not  
313 the observable state or mass eigenstate. The object describing this mixing in the quark sector is the  
314 Cabibbo-Kobayashi-Maskawa (CKM) matrix [13].

315 In the leptonic sector, the masses of the neutrino cause mismatch between the mass and the flavor  
316 eigenstate allowing for neutrino oscillations. This topic will be addressed in further details in Section  
317 1.2.2.

### 318 1.1.2 Limits of the standard model

319 The SM has been successful at describing many of phenomena observed in experiments. However,  
320 some questions remain unanswered, among which:

- 321 — Dark matter and dark energy. Cosmological observations – such as the acceleration of the  
322 expansion of the universe and the rotational speed of galaxies, for example – indicate the  
323 presence of unknown energy and matter in the universe. The  $\Lambda$ CDM model [14, 15] indicates  
324 that only 4.5% of the total energy in the universe is described by the SM. The supplementary  
325 mass – dark matter – accounts for 22.5% of the missing energy and the rest is dark energy.
- 326 — Baryonic Asymmetry of the universe. The universe is mainly made of matter. The Sakharov  
327 theory [16] to explain the deficit of antimatter require the breaking of the C and CP symmetries.  
328 The CP violation is allowed in the strong sector, but not to the magnitude necessary to explain  
329 the quasi absence of baryonic anti-matter. Other mechanism must exist to explain this deficit.
- 330 — Fermion masses. The large mass difference between the fermions is not explained by the SM.
- 331 — The SM includes 26 numerical parameters, the values of which are determined only through  
332 experimental measurements. At least 20 of these parameters are related to flavor physics. In  
333 electroweak theory nothing dictates the values of the interaction couplings and masses. This  
334 reflects a deeper limitation of the model, as it provides no theoretical explanation for the values  
335 of these parameters, which suggests that a more fundamental theory may be required.
- 336 — Strong CP problem. Theoretically it is possible to have violation of CP symmetry in strong  
337 interactions. Experimentally, however, no such asymmetry has been found, implying that the  
338 coefficient of this term is very close to zero. This fine tuning is considered unnatural.
- 339 — Non-unification of couplings. The gauge couplings of the SU(3), SU(2) and U(1) groups are  
340 independent quantities. Due to higher-order corrections, each of these is actually a function  
341 of the typical energy scale  $Q$  relevant to the process. In many grand unified theories the  
342 three gauge couplings are predicted to meet at some high energy unification. However, this  
343 unification does not occur when the couplings are extrapolated using the SM model expression.
- 344 — Gravitation. The SM do not include the Gravitational interaction and is incompatible with the  
345 general relativity.

## 346 1.2 The Neutrinos

347 As introduced in the previous section, the neutrino are the neutral leptons of the Standard Model  
348 (SM). They were first theorized by Wolfgang Ernst Pauli in 1978 [17] to solve the problem of the  
349  $\beta$ -decay continuous spectrum. Indeed if the  $\beta$ -decay was a two body reaction  ${}^A_Z X \rightarrow {}^A_{Z+1} Y + e^-$ , the  
350 conservation of momentum would force the charged lepton to be mono-energetic, but the measured  
351 spectrum was continuous. To solve this problem Pauli theorized the emission of a neutral particle  
352  ${}^A_Z X \rightarrow {}^A_{Z+1} Y + e^\pm + \nu$ , the neutrino. This particle had to be light, neutral, and interact weakly with  
353 matter.



FIGURE 1.2 – Feynman diagrams of the charged current (on the left) and the neutral current (on the right) for a lepton  $l$  and its corresponding neutrino  $\nu_l$ .

354 We must wait 1956 for a collaboration led by Frederick Reines and Clyde Cowan for the first obser-  
 355 vation of the neutrino [18, 19] via the Inverse Beta Decay (IBD) reaction

$$\bar{\nu} + p \rightarrow e^+ + n \quad (1.1)$$

356 Following this discovery, numerous experiments were setup to study its properties. Some of the  
 357 notable discoveries include the discovery in 1962, by a collaboration led by Leon Lederman, Melvin  
 358 Schwartz and Jack Steinberg, of the muon neutrino flavor [20].

359 Soon after, the Homestake experiment, which was measuring the neutrino produced by the proton-  
 360 proton fusion cycle in the sun, reported a deficit of factor  $\sim 3$  [21] in comparison to the Standard  
 361 Solar Model predictions. This anomaly, referred as the *solar neutrino problem* remained unexplained  
 362 until the neutrino oscillation was theorized and proven. Bruno Pontecorvo first suggested a  $\nu \leftrightarrow \bar{\nu}$   
 363 oscillation [22], later revisited by Maki et al. to a two flavor oscillation  $\nu_e \leftrightarrow \nu_\mu$  [23]. The discovery of  
 364 the  $\tau$  lepton 1976 [24] and its associated neutrino  $\nu_\tau$  [25] led to the extension to three flavor oscillation.

365 This three flavor oscillation was confirmed by the observation of the  $\nu_\mu \leftrightarrow \nu_\tau$  oscillation [26] in 1998  
 366 by the Super-Kamiokande experiment.

### 367 1.2.1 Coupling and interactions

368 The SM, as originally defined, contains no right-handed neutrino (right helicity) since only left-  
 369 handed neutrinos have been observed [27], implying that the neutrinos are massless. Neutrinos  
 370 actually do have a very small mass, with the current best limit  $m_\nu < 0.45$  eV at 90% confidence  
 371 level, obtained from the measurement of the electron energy spectrum in tritium beta decay by the  
 372 KATRIN experiment [28]. They only couple – interact – through the  $W^\pm$  and  $Z^0$  bosons. The coupling  
 373 with a  $W^\pm$  boson is the *charged current*, a charge is exchanged via the  $W$  boson, and coupling with  $Z^0$   
 374 is the neutral current, no charge is exchanged. The Feynman diagrams representing those interaction  
 375 are presented in figure 1.2.

376 As explained in Section 1.1, those interactions preserve the leptonic quantum number  $L$ . In the  
 377 absence of neutrino mass, the leptonic flavour numbers  $L_e$ ,  $L_\mu$  and  $L_\tau$  are also exactly conserved.  
 378 However, the existence of neutrino masses allow for lepton flavor violating transition such as the  
 379 oscillation  $\nu_\alpha \rightarrow \nu_{\beta \neq \alpha}$  but also process such as  $\mu^+ \rightarrow e^+ + \gamma$  or  $\mu^+ \rightarrow e^+ e^+ e^-$ . The latter that are  
 380 heavily suppressed – their probability of occurring is extremely low in comparison to other processes  
 381 – in the absence of new physics [29].

### 382 1.2.2 Oscillation

383 Neutrino oscillations occur due to the fact that the flavor states, in which neutrinos are produced and  
 384 detected, are quantum superpositions of mass eigenstates. This results in neutrinos changing their  
 385 flavor as they propagate, with the oscillation parameters depending on the differences in the squared  
 386 masses of the mass eigenstates. More strictly speaking, their mass induces a mismatch between the  
 387 *flavor states*  $|\nu_e\rangle$ ,  $|\nu_\mu\rangle$  and  $|\nu_\tau\rangle$  which are the state in which the particle interacts – the states in the  
 388 diagrams in Figure 1.2 – and the *mass states*  $|\nu_1\rangle$ ,  $|\nu_2\rangle$  and  $|\nu_3\rangle$  which hold the momentum and mass  
 389 of the particle.

390 Thus the flavor state  $|\nu_\alpha\rangle$  is a quantum superposition of mass eigenstates, and can be written

$$|\nu_\alpha\rangle = \sum_{i=1}^3 U_{\alpha,i} |\nu_i\rangle \quad (1.2)$$

391 and reciprocally

$$|\nu_i\rangle = \sum_{\alpha \in e, \mu, \tau} U_{\alpha,i}^* |\nu_\alpha\rangle \quad (1.3)$$

392 where  $i$  indexes the mass states,  $\alpha$  the flavor states and the  $U_{\alpha,i}$  are the mixing elements that governs  
 393 the probability amplitudes for the neutrino flavor transitions. In the three-families framework, this  
 394 mixing is represented by the  $3 \times 3$  Pontecorvo-Maki-Nakagawa-Sakata matrix [23]  $U_{\text{PMNS}}$

$$\begin{pmatrix} |\nu_e\rangle \\ |\nu_\mu\rangle \\ |\nu_\tau\rangle \end{pmatrix} = U_{\text{PMNS}} \begin{pmatrix} |\nu_1\rangle \\ |\nu_2\rangle \\ |\nu_3\rangle \end{pmatrix} = \begin{pmatrix} U_{e1} & U_{e2} & U_{e3} \\ U_{\mu 1} & U_{\mu 2} & U_{\mu 3} \\ U_{\tau 1} & U_{\tau 2} & U_{\tau 3} \end{pmatrix} \begin{pmatrix} |\nu_1\rangle \\ |\nu_2\rangle \\ |\nu_3\rangle \end{pmatrix} \quad (1.4)$$

395 This matrix is considered to be unitary but this property still needs to be corroborated [30]. Now,  
 396 considering a neutrino produced as  $|\nu_\alpha\rangle$  that propagates over a distance  $x$  during a time  $t$ , the  
 397 Schrödinger equation [31] can be written as:

$$|\nu_\alpha(x, t)\rangle = \sum_{i=1}^3 U_{\alpha,i} e^{-i(E_i t - p_i x)} |\nu_i\rangle \quad (1.5)$$

398 where  $E_i$  and  $p_i$  stand for the energy and momentum of the neutrino mass states respectively. By  
 399 going back from the mass space to the flavor space, Eq. 1.5 becomes:

$$|\nu_\alpha(x, t)\rangle = \sum_{\beta \in e, \mu, \tau} U_{\beta,i}^* \left( \sum_{i=1}^3 U_{\alpha,i} e^{-i(E_i t - p_i x)} \right) |\nu_\beta\rangle \quad (1.6)$$

400 A neutrino created as  $|\nu_\alpha\rangle$  thus propagates as the linear superposition of the three flavor states. Because  
 401 the mass of the neutrino is extremely small, we can consider that they are ultra-relativistic ( $E \sim p \gg m$ ). Using natural units ( $c = \hbar = 1$ ):

$$E_i = \sqrt{p^2 + m_i^2} \simeq p + \frac{m_i^2}{2p} \simeq E + \frac{m_i^2}{2E} \quad (1.7)$$

<sup>403</sup> then the probability to observe a neutrino produced in state  $|\nu_\alpha\rangle$  in a state  $|\nu_\beta\rangle$  can be written<sup>1</sup>:

$$P_{\nu_\alpha \rightarrow \nu_\beta} = |\langle \nu_\beta | \nu_\alpha \rangle|^2 = \sum_{i,j=1}^3 U_{\alpha,i}^* U_{\beta,i} U_{\alpha,j}^* U_{\beta,j} e^{-i \frac{\Delta m_{ji}^2 L}{2E}} \quad (1.8)$$

<sup>404</sup> where  $L = ct$  is the propagation distance of the neutrino,  $E$  is the neutrino energy and  $\Delta m_{ji}^2 =$   
<sup>405</sup>  $m_j^2 - m_i^2$  is the *mass splitting*, the difference between the square of the eigenvalues of two mass states.

<sup>406</sup> Under unitary assumptions, the PMNS matrix can also also be decomposed in three rotational ma-  
<sup>407</sup> trices:

$$U_{\text{PMNS}} = \begin{pmatrix} 1 & 0 & 0 \\ 0 & \cos \theta_{23} & \sin \theta_{23} \\ 0 & -\sin \theta_{23} & \cos \theta_{23} \end{pmatrix} \begin{pmatrix} \cos \theta_{13} & 0 & \sin \theta_{13} e^{-i\delta_{\text{CP}}} \\ 0 & 1 & 0 \\ -\sin \theta_{13} e^{i\delta_{\text{CP}}} & 0 & \cos \theta_{13} \end{pmatrix} \begin{pmatrix} \cos \theta_{12} & \sin \theta_{12} & 0 \\ -\sin \theta_{12} & \cos \theta_{12} & 0 \\ 0 & 0 & 1 \end{pmatrix} \quad (1.9)$$

<sup>408</sup> where the parameters  $\theta_{12}$ ,  $\theta_{23}$ ,  $\theta_{13}$  are the *mixing angles*. The parameter  $\delta_{\text{CP}}$  is a CP violation phase  
<sup>409</sup> that quantifies the matter-antimatter asymmetry in the leptonic sector. The parameters  $\theta_{12}$  and  $\Delta m_{21}^2$   
<sup>410</sup> are commonly attributed to a so-called *solar sector*, while the parameters  $\theta_{13}$  and  $\Delta m_{31}^2$  belong to the  
<sup>411</sup> *reactor sector* and  $\theta_{23}$  and  $\Delta m_{32}^2$  the *atmospheric sector*. The neutrino oscillation is thus characterized  
<sup>412</sup> by 7 parameters: the three mixing angles ( $\theta_{12}, \theta_{13}, \theta_{23}$ ), the three mass splitting ( $\Delta m_{21}^2, \Delta m_{31}^2, \Delta m_{32}^2$ )  
<sup>413</sup> and the CP violation phase  $\delta$ . These three mass splittings are constrained by the relation

$$\Delta m_{21}^2 + \Delta m_{32}^2 - \Delta m_{31}^2 = 0 \quad (1.10)$$

<sup>414</sup> The neutrinos interact weakly with matter. But even so, the travel through dense matter, such as the  
<sup>415</sup> Earth's crust, can impact their propagation probability. These *matter effects* were introduced for the  
<sup>416</sup> first time by Lincoln Wolfenstein, Stanislas Mikheyev and Alexei Smirnov in 1978 [33]. They result  
<sup>417</sup> from forward elastics scattering of neutrinos with the medium (the momentum of the neutrino is  
<sup>418</sup> unchanged). The charged and neutral current Feynman diagrams are presented in Figure 1.3. This  
<sup>419</sup> results in a supplementary potential in the Hamiltonian, impacting the oscillation probability.

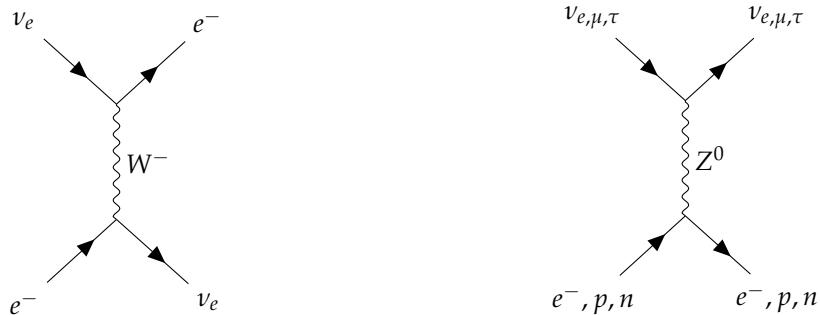


FIGURE 1.3 – Feynman diagrams of the of charged current matter effect (on the left) and the neutral current matter effect (on the right). Only the electronic neutrino is sensitive to charged current, whereas every neutrinos are sensitive to neutral current.

1. Actually Eq. 1.7 and 1.8 make a few more assumptions, such as the fact that every mass state has the same momentum. “Paradoxes of Neutrino Oscillations” from Akhmedov and Smirnov [32] go through them and demonstrate the validity of the method presented in this chapter.

420 **1.2.3 Phenomenology**

421 The neutrino experiments can be divided into two main categories: the disappearance experiments,  
 422 which observe a deficit of a specific flavor of neutrinos in the detector compared with the expected  
 423 source flux, and the appearance experiments that search for an excess of a flavor. By placing them at  
 424 different distances – baselines – we can favor the appearance or disappearance of different neutrino  
 425 flavors. As an illustration of the effect of the baseline, the survival probability of  $\bar{\nu}_e$  with respect of  
 426 the baseline is presented in Figure 1.4.

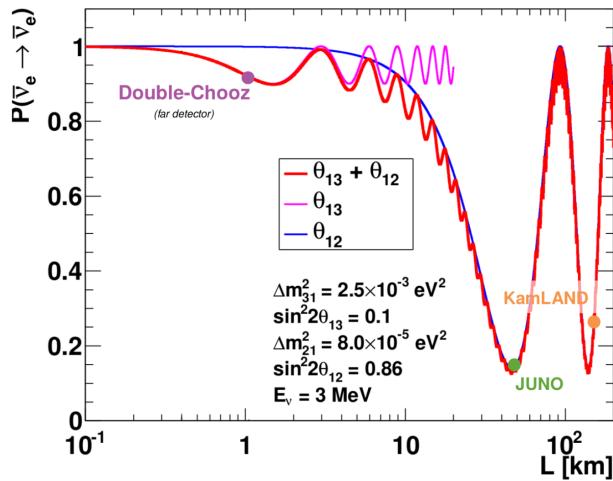


FIGURE 1.4 – Survival probability of  $\bar{\nu}_e$  as a function of the baseline. The energy of the neutrinos is 3 MeV. The baseline of Double-Chooz, JUNO and KamLAND are reported. Figure taken from Ref. [34].

427 The experiments are also characterized by the neutrino sources they observe. For neutrino studies,  
 428 the common sources are solar neutrinos, atmospheric neutrino produced from the interactions of  
 429 high energy particles in the upper atmosphere, accelerators by firing particle beams on a target, and  
 430 nuclear power plant reactors.

431 One of the major questions right now is the Mass Ordering (MO), the sign of the mass splitting terms  
 432  $\Delta m_{31}^2$ , and thus  $\Delta m_{32}^2$ , is not known. We consider two hypotheses: Normal Ordering (NO) where  
 433  $m_1 < m_2 < m_3$  and the Inverted Ordering (IO) where  $m_3 < m_1 < m_2$ . An illustration of the MO is  
 434 presented in Figure 1.5. This topic will be further discussed in Section 1.2.4.

435 **Solar sector ( $\theta_{12}$ ,  $\Delta m_{21}^2$ )**

436 The measurement of the solar sector parameters  $\theta_{12}$  and  $\Delta m_{21}^2$  has been done in two different ways.  
 437 From the measurements of the solar neutrino flux in experiments like Super Kamiokande [35] and  
 438 by extracting the parameter from the reactor  $\bar{\nu}_e$  spectrum, as done by the KamLand-Zen experiment  
 439 [36, 37]. Those results are further constrained by measurements of short-baseline experiments and  
 440 accelerator data. The Particle Data Group (PDG) in its latest edition [1] reports the value from global  
 441 fit efforts from [35, 37]

$$\sin^2 \theta_{12} = 0.307^{+0.013}_{-0.012}$$

$$\Delta m_{21}^2 = 7.53 \pm 0.18 \cdot 10^{-5} \text{ eV}^2$$

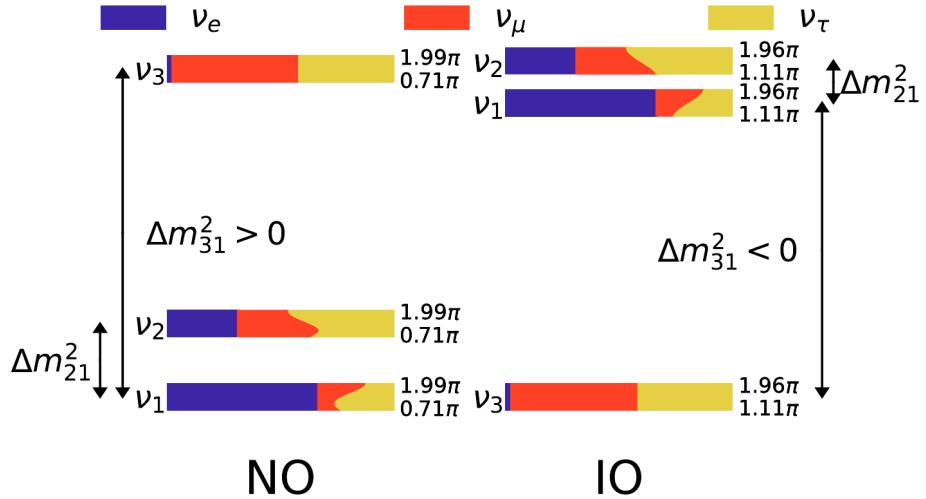


FIGURE 1.5 – Illustration of the Mass Ordering, on the left the Normal Ordering and on the right the Inverted Ordering

442 The intervals are 68% confidence level. The CPT invariance is assumed.

443 **Reactor sector ( $\theta_{13}$ )**

444 Direct measurements of  $\theta_{13}$  are currently derived from the reactor  $\bar{\nu}_e$  disappearance at short baseline  
445  $L \sim 1$  km. Alternatively, limits can also be obtained from solar neutrino data or accelerator based  
446  $\nu_\mu \rightarrow \nu_e$  experiments. The PDG reports its best value as the average of the results T2K [38], DayaBay  
447 [39, 40], Double-Chooz [41] and Reno [42, 43]

$$\sin^2 \theta_{13} = 2.19 \pm 0.07$$

448 The intervals are 68% confidence level. The CPT invariance is assumed.

449 **Atmospheric sector ( $\theta_{23}, \Delta m_{32}^2$ ) and  $\delta_{CP}$**

450 The parameters  $\theta_{23}$  and  $\Delta m_{32}^2$  are currently constrained by the measurements the relative fluxes of  
451 atmospheric neutrino flavors and by accelerator-based experiments. The mass splitting term can also  
452 be extracted from the  $\bar{\nu}_e$  spectrum at short baselines. The PDG uses the results of IceCube [44], T2K  
453 [38], NOvA [45], MINOS [46] and Super Kamiokande (SK) [47] for  $\theta_{13}$ . For  $\Delta m_{31}^2$ , they also use the  
454 data from Daya Bay [39] and RENO [43].

455 **Assuming normal ordering**

$$\sin^2 \theta_{23} = 0.553^{+0.016}_{-0.024}$$

$$\Delta m_{32}^2 = 2.445 \pm 0.028 \cdot 10^{-3} \text{ eV}^2$$

**Assuming inverted ordering**

$$\sin^2 \theta_{23} = 0.558^{+0.015}_{-0.021}$$

$$\Delta m_{32}^2 = -2.529 \pm 0.029 \cdot 10^{-3} \text{ eV}^2$$

<sup>456</sup> The CP violation phase  $\delta_{\text{CP}}$  is measured from the  $\nu_e$  appearance in atmospheric and accelerator  
<sup>457</sup> experiments. PDG uses the data from T2K [38], NOVA [45] and Super Kamiokande [47]. The  
<sup>458</sup> reported value is in  $\pi$  radians with  $0 < \delta_{\text{CP}} < 2\pi$ . This value of  $\delta_{\text{CP}}$  assumes Normal Ordering.

$$\delta_{\text{CP}} = 1.19 \pm 0.22$$

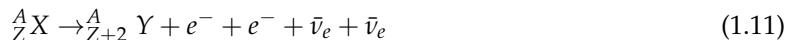
<sup>459</sup> The intervals are 68% confidence level. The CPT invariance is assumed.

#### <sup>460</sup> 1.2.4 Open questions

<sup>461</sup> Neutrino experiments have continued to produce more and more precise and refined results over  
<sup>462</sup> the last 25 years. However, some questions remain unanswered and will be at the center of attention  
<sup>463</sup> of the neutrino physics community for the years to come. Some of them are reviewed in this section.

##### <sup>464</sup> Nature of the neutrino

<sup>465</sup> The particles in the SM are Dirac particles (i.e. there is a distinction between particles and antiparticles). The modern model of neutrino interaction and oscillation has been developed under this  
<sup>466</sup> postulate. However, Ettore Majorana in 1937 formulated another way to introduce mass terms in  
<sup>467</sup> the SM Lagrangian [48]. The condition for this mass term is only applicable to the neutrino. If the  
<sup>468</sup> neutrino is a Majorana particle, meaning it is its own antiparticle, lepton number violation would  
<sup>469</sup> occur during certain processes such as neutrinoless double beta decay. This violation is predicted  
<sup>470</sup> because in Majorana interactions, neutrinos can annihilate themselves, violating lepton number  
<sup>471</sup> conservation by two units. To determine the nature of the neutrino (Majorana or Dirac nature),  
<sup>472</sup> physicist observe the rare phenomenon of  $\beta\beta$  decay – double beta decay:



<sup>474</sup> In this scenario, if the neutrino are Majorana particles, one could observe a theoretical rarer event  
<sup>475</sup>  $0\nu\beta\beta$  decay:



<sup>476</sup> where the two neutrinos annihilate in the interaction. The signature would be the presence of quasi-  
<sup>477</sup> monoenergetic events at the higher end of the  $\beta\beta$  decay energy spectrum.

<sup>478</sup> The implication of the Majorana nature of the neutrino implies, among other:

- <sup>479</sup> — The lepton number violation.
- <sup>480</sup> — Possible explanation of small neutrino masses via seesaw mechanism [49].
- <sup>481</sup> — Open the possibility of generating the baryon asymmetry via leptogenesis.

<sup>482</sup> As of today, the nature of the neutrino hasn't been established.

##### <sup>483</sup> Absolute mass of the neutrino

<sup>484</sup> Studies of neutrino oscillation allow us to measure the mass splitting terms  $\Delta m_{ij}^2 = m_i^2 - m_j^2$  but  
<sup>485</sup> do not provide informations about the absolute mass scale of the neutrinos. The most stringent  
<sup>486</sup> upper limits on the absolute neutrino mass come from the KATRIN experiment, which measures  
<sup>487</sup> the electron energy spectrum in tritium beta decay. By precisely studying the tail of this spectrum,  
<sup>488</sup> where small energy differences due to the neutrino mass can manifest, KATRIN is able to constrain  
<sup>489</sup> the neutrino mass with a high degree of sensitivity. The latest result gives an upper limit at a 90%  
<sup>490</sup> confidence level of  $m_\nu < 0.45$  eV [28].

491 **Neutrino Mass Ordering (NMO)**

492 As introduced in Section 1.2.3, current experiments are only sensitive to  $|\Delta m_{32}^2|$ , blinded to the sign of  
 493 the mass split. We are thus unable to differentiate between the Normal Ordering (NO)  $m_1 < m_2 < m_3$   
 494 and the Inverted Ordering (IO)  $m_3 < m_1 < m_2$ . The nature of the NMO has important implications:

- 495 — It will help in the determination of the lower limit of the mass scale.
- 496 — The nature of the neutrino mass ordering has significant implications for various experiments,  
 497 particularly neutrinoless double beta decay ( $0\nu\beta\beta$ ). If the mass ordering is inverted, the ef-  
 498 fective mass for  $0\nu\beta\beta$  is likely to be higher, which would increase the decay's detectability.  
 499 Conversely, a normal ordering would imply a lower effective mass, making detection more  
 500 challenging.
- 501 — The current  $\delta_{CP}$  measurements have different minima depending on the mass ordering as  
 502 illustrated in Figure 1.6.
- 503 — More broadly, knowledge of the mass ordering and the neutrino mass scale have impacts on  
 504 the generation of lepton in early universe and cosmology [50].

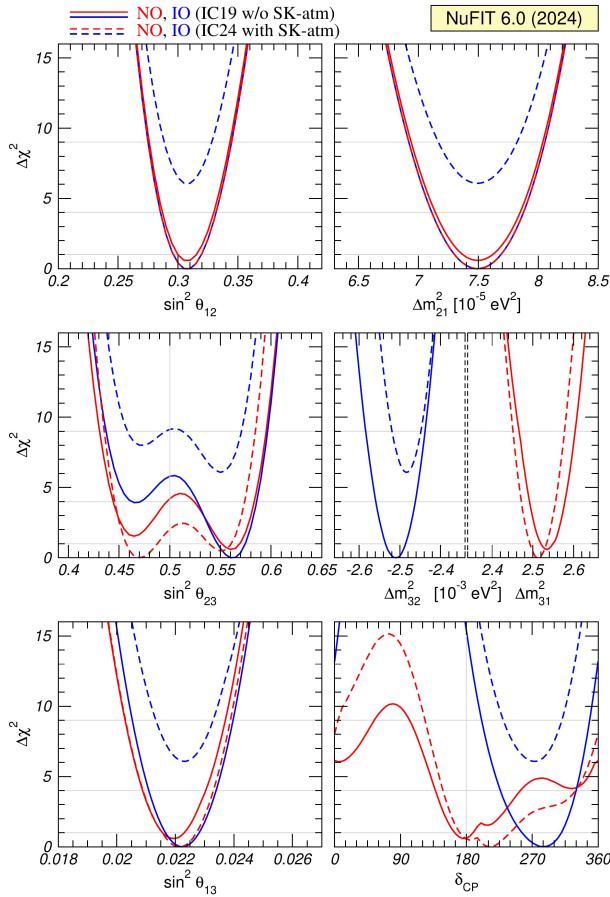


FIGURE 1.6 – Global  $3\nu$  oscillation analysis. The red (blue) curves are for Normal (Inverted) Ordering. The solid (dashed) lines are obtained without (with) the inclusion of the tabulated atmospheric  $\chi^2$  data from SK and IC. As atmospheric mass-squared splitting we use  $\Delta m_{31}^2$  for NO and  $\Delta m_{32}^2$  for IO. Analysis and plots from the NuFit analysis group [51, 52].

505 The determination of the mass ordering will probably be solved in the next decade. Multiple ex-  
 506 periments under construction have this topic as part of their main physics program, notably the

507 Jiangmen Underground Neutrino Observatory (JUNO) [53] – the subject of this thesis – via the  
 508 measurement of the  $\bar{\nu}_e$  spectrum from nuclear reactors; the long-baseline DUNE experiment [54]  
 509 which will detect neutrinos from the Fermilab accelerator; KM3Net/ORCA and Hyper-Kamiokande  
 510 through the measurements of the atmospheric neutrino flux [55, 56].

511 **Octant of  $\theta_{23}$**

512 The latest results from the NuFit analysis group [51] show two minima for  $\theta_{13}$  (Fig. 1.6), one in the  
 513 lower octant  $\theta_{23} < \pi/4$  and one in the upper octant  $\theta_{23} > \pi/4$ . This octant has implications for the  
 514 oscillation and neutrino mass theories. It will be measured by future neutrino experiments such as  
 515 DUNE and HK.

516 **Breaking of CP symmetry in the leptonic sector**

517 The CP symmetry is known to be broken in the baryonic sector [6] however its violation in the  
 518 leptonic sector is still unknown. The CP violation exists if the PMNS matrix possesses an imaginary  
 519 part, i.e. if  $\text{Im}(e^{i\delta_{CP}}) \neq 0$ . The latest measurements (Fig. 1.6) do not give any certainties about the  
 520 value of  $\delta_{CP}$ ; it will be determined by future experiments such as DUNE and HK. The violation of CP  
 521 in the leptonic sector has strong implications for the matter-antimatter asymmetry in the universe.

522 **Sterile neutrino**

523 Sterile neutrinos are hypothetical particles that extend the Standard Model by introducing a fourth,  
 524 non-interacting neutrino family. They have been proposed to explain certain anomalies observed  
 525 in short-baseline neutrino experiments, such as the LSND [57] and MiniBooNE [58] experiments,  
 526 where observed oscillations cannot be accounted for by the known three neutrino flavors. This  
 527 sterile neutrino would not interact with the other particles of the SM except via gravity and could be  
 528 observed only via the oscillations of  $\nu_e, \nu_\mu, \nu_\tau \rightarrow \nu_s$  where  $\nu_s$  would be this hypothetical fourth sterile  
 529 family. This theory has been proposed multiple times to explain anomalies in the oscillations data,  
 530 but for now no studies have concluded on the existence of the sterile neutrino. One of the possible  
 531 probes for this fourth family would be the measurement of the non-unitarity of the PMNS matrix,  
 532 meaning that  $\nu_e$  is a linear combination of  $\nu_1, \nu_2, \nu_3$  and a theoretical  $\nu_4$ . This unitarity test can be  
 533 done via precise measurements of the mixing angles.

<sup>534</sup> **Chapter 2**

<sup>535</sup> **The JUNO experiment**

<sup>536</sup>

*"Ave Juno, rosae rosam, et spiritus rex". It means nothing but I found it in tone.*

<sup>537</sup>

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<sup>561</sup>

<sup>562</sup>

<sup>563</sup>

<sup>564</sup>

The first idea of a medium baseline ( $\sim 52$  km) experiment, was explored in 2008 [59] where it was demonstrated that the Neutrino Mass Ordering (NMO) could be determined by a medium baseline experiment if  $\sin^2(2\theta_{13}) > 0.005$  without the requirements of accurate knowledge of the reactor antineutrino spectra and the value of  $\Delta m_{32}^2$ . From this idea is born the Jiangmen Underground Neutrino Observatory (JUNO) experiment.

<sup>570</sup>

<sup>571</sup>

<sup>572</sup>

<sup>573</sup>

JUNO is a neutrino detection experiment under construction located in China, in Guangdong province, near the city of Kaiping. Its main objectives are the determination of the mass ordering at the  $3\text{-}4\sigma$  level in 6 years of data taking and the measurement at the sub-percent precision of the oscillation parameters  $\Delta m_{21}^2$ ,  $\sin^2 \theta_{12}$ ,  $\Delta m_{32}^2$  and with less precision  $\sin^2 \theta_{13}$ [60].



FIGURE 2.1 – **On the left:** Location of the JUNO experiment and its reactor sources in southern china. **On the right:** Aerial view of the experimental site

For this JUNO will measure the electronic anti-neutrinos ( $\bar{\nu}_e$ ) flux coming from the nuclear reactors of Taishan, Yangjiang, for a total power of 26.6 GW<sub>th</sub>, and the Daya Bay power plant to a lesser extent. All of those cores are the second-generation pressurized water reactors CPR1000, which is a derivative of Framatome M310. Details about the power plants characteristics and their expected flux of  $\bar{\nu}_e$  can be found in the table 2.1. The distance of 53 km has been specifically chosen to maximize the disappearance probability of the  $\bar{\nu}_e$ . The data taking is scheduled to start early 2025.

## 2.1 Reactor Neutrinos physics in JUNO

JUNO will try to determine the NMO and to bring at the few per mille level our knowledge of  $\Delta m_{31}^2$ ,  $\Delta m_{21}^2$  and  $\sin^2(2\theta_{12})$  via the precision analysis of the spectrum of the visible energy left by reactor antineutrinos in its detector.

### 2.1.1 Antineutrino spectrum measured in JUNO

To some extent, this analysis is equivalent to extracting from this spectrum the oscillation probability [60] :

$$P(\bar{\nu}_e \rightarrow \bar{\nu}_e) = 1 - \sin^2 2\theta_{12} c_{13}^4 \sin^2 \frac{\Delta m_{21}^2 L}{4E} - \sin^2 2\theta_{13} \left[ c_{12}^2 \sin^2 \frac{\Delta m_{31}^2 L}{4E} + s_{12}^2 \sin^2 \frac{\Delta m_{32}^2 L}{4E} \right]$$

Where  $s_{ij} = \sin \theta_{ij}$ ,  $c_{ij} = \cos \theta_{ij}$ ,  $E$  is the  $\bar{\nu}_e$  energy and  $L$  is the baseline. We can see the sensitivity to the NMO in the dependency to  $\Delta m_{32}^2$  and  $\Delta m_{31}^2$  causing a phase shift of the spectrum as we can see in the Figure 2.2.

In practice, a fit to the grey distribution of Figure 2.3 will be performed. It is the sum of two components :signal (black) and bacgrounds (colored). Reactor antineutrinos are detected by JUNO via Inverse Beta Decays (IBD) :  $\bar{\nu}_e + p \rightarrow e^+ + n$ . The energy spectrum under investigation is therefore that of the reconstructed  $e^+$  visible energy. The black signal spectrum is therefore the sum of the antineutrino differential fluxes from all reactors and reaching the detecteur, weighted by the oscillation probability of Eq 2.1.1 and the IBD differential cross section and convoluted with detection effects. These various ingredients are theoretically modelled in order to provide the probability density function (PDF) to be used in the fit.

To reach JUNO's goals, it takes that this experimental spectrum still bears sizeable traces of the very small phase shift mentioned above. Most notably, the following requirements must be fulfilled :

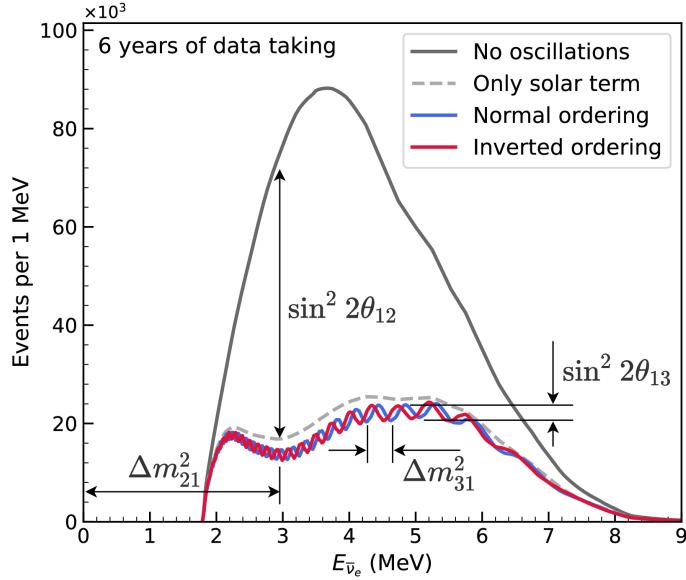


FIGURE 2.2 – Expected number of neutrinos event per MeV in JUNO after 6 years of data taking. The black curve shows the flux if there was no oscillation. The light gray curve shows the oscillation if only the solar terms are taken in account ( $\theta_{12}$ ,  $\Delta m_{21}^2$ ). The blue and red curve shows the spectrum in the case of, respectively, NO and IO. The dependency of the oscillation to the different parameters are schematized by the double sided arrows. We can see the NMO sensitivity by looking at the fine phase shift between the red and the blue curve.

- 598 1. An energy resolution of  $3\%/\sqrt{E(\text{MeV})}$  to be able to distinguish the fine structure of the fast  
599 oscillation.
- 600 2. An energy scale known at the better than the 1% level.
- 601 3. A baseline between 40 and 65 km to maximise the  $\bar{\nu}_e$  oscillation probability. The optimal  
602 baseline would be 58 km and JUNO baseline is 53 km.
- 603 4. At least  $\approx 100,000$  events. This is the necessary statistics to reach JUNO's canonical sensitivity  
604 after 6 years of data taking.

#### 605 $\bar{\nu}_e$ flux coming from nuclear power plants

606 To get such high measurements precision, it is necessary to have a very good understanding of the  
607 sources characteristics. For its NMO and precise measurement studies, JUNO will observe the energy  
608 spectrum of neutrinos coming from the nuclear power plants Taishan and Yangjiang's cores, located  
609 at 53 km of the detector to maximise the disappearance probability of the  $\bar{\nu}_e$ .

610 The  $\bar{\nu}_e$  coming from reactors are emitted from  $\beta$ -decay of unstable fission fragments. The Taishan  
611 and Yangjiang reactors are Pressurised Water Reactor (PWR), the same type as Daya Bay. In those  
612 type of reactor more the 99.7 % and  $\bar{\nu}_e$  are produced by the fissions of four fuel isotopes  $^{235}\text{U}$ ,  $^{238}\text{U}$ ,  
613  $^{239}\text{Pu}$  and  $^{241}\text{Pu}$ . The neutrino flux per fission of each isotope is determined by the inversion of the  
614 measured  $\beta$  spectra of fission product [62–66] or by calculation using the nuclear databases [67, 68].

615 The neutrino flux coming from a reactor at a time  $t$  can be predicted using

$$\phi(E_\nu, t)_r = \frac{W_{th}(t)}{\sum_i f_i(t) e_i} \sum_i f_i(t) S_i(E_\nu) \quad (2.1)$$

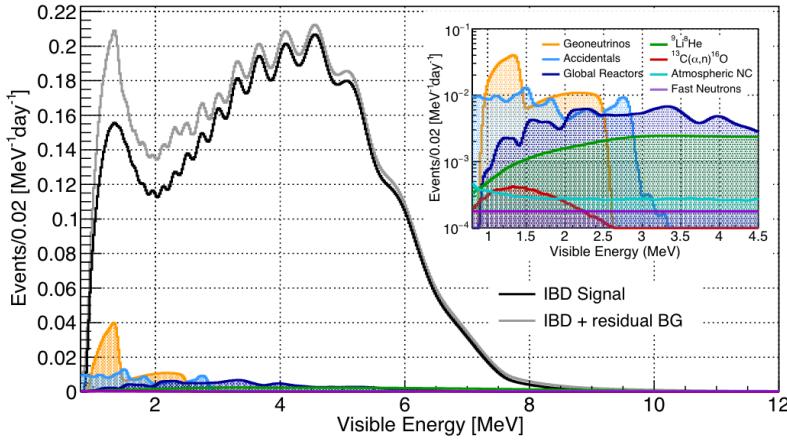


FIGURE 2.3 – Expected visible energy spectrum measured with the LPMT system with (grey) and without (black) backgrounds. The background amount for about 7% of the IBD candidate and are mostly localized below 3 MeV [61]

Reactor	Power (GW <sub>th</sub> )	Baseline (km)
Taishan	9.2	52.71
Core 1	4.6	52.77
Core 2	4.6	52.64
Yangjiang	17.4	52.46
Core 1	2.9	52.74
Core 2	2.9	52.82
Core 3	2.9	52.41
Core 4	2.9	52.49
Core 5	2.9	52.11
Core 6	2.9	52.19
Daya Bay	17.4	215
Huizhou	17.4	265

TABLE 2.1 – Characteristics of the nuclear power plants observed by JUNO.

616 where  $W_{th}(t)$  is the thermal power of the reactor,  $f_i(t)$  is the fraction fission of the  $i$ th isotope,  $e_i$  its  
 617 thermal energy released in each fission and  $S_i(e_\nu)$  the neutrino flux per fission for this isotope.

618 The latter flux is difficult to predict. To evaluate JUNO’s sensitivity and to serve as a starting point  
 619 in the spectrum PDF, the Huber-Mueller model is used [63], corrected using Daya Bays data [69] to  
 620 account for a  $\sim 5\%$  deficit with respect to models, referred to as the reactor antineutrino anomaly  
 621 [70], and for a discrepancy between models and data in the spectral shape (the so call 5 MeV bump).

622 In addition to those prediction, a satellite experiment named TAO[71] will be setup near the reactor  
 623 core Taishan-1 to measure with an energy resolution of 2% at 1 MeV the neutrino flux coming from  
 624 the core, more details can be found in Section 2.5.1. It will help identifying unknown fine structure  
 625 and give more insight on the  $\bar{\nu}_e$  flux coming from this reactor.

### 626 2.1.2 Background spectra

627 Considering the close reactor neutrinos flux as the main signal, the signals that are considered as  
 628 background are:

- 629 — The geoneutrinos producing background in the  $0.511 \sim 2.7$  MeV region.

630 — The neutrinos coming from the other nuclear reactors around Earth.

631 In addition to all those physics signal, non-neutrinos signal that would mimic an IBD will also be  
632 present. It is composed of:

633 — The signal coming from radioactive decay ( $\alpha$ ,  $\gamma$ ,  $\beta$ ) from natural radioactive isotopes in the  
634 material of the detector.

635 — Cosmogenic event such as fast neutrons and activated isotopes induced by muons passing  
636 through the detector, most notably the spallation on  $^{12}\text{C}$ .

637 All those events represent a non-negligable part of the spectrum as shown in Figure 2.3.

## 638 2.2 Other physics

639 While the design of JUNO is tailored to measure  $\bar{\nu}_e$  coming from nuclear reactor, JUNO will be able  
640 to detect neutrinos coming from other sources thus allowing for a wide range of physics studies as  
641 detailed in the table 2.2 and in the following sub-sections.

Research	Expected signal	Energy region	Major backgrounds
Reactor antineutrino	60 IBDs/day	0-12 MeV	Radioactivity, cosmic muon
Supernova burst	5000 IBDs at 10 kpc	0-80 MeV	Negligible
DSNB (w/o PSD)	2300 elastic scattering		
Solar neutrino	2-4 IBDs/year	10-40 MeV	Atmospheric $\nu$
Atmospheric neutrino	hundreds per year for $^{8}\text{B}$	0-16 MeV	Radioactivity
Geoneutrino	hundreds per year	0.1-100 GeV	Negligible
	$\approx 400$ per year	0-3 MeV	Reactor $\nu$

TABLE 2.2 – Detectable neutrino signal in JUNO and the expected signal rates and major background sources

### 642 Geoneutrinos

643 Geoneutrinos designate the antineutrinos coming from the decay of long-lived radioactive elements  
644 inside the Earth. The 1.8 MeV threshold necessary for the IBD makes it possible to measure geoneu-  
645 trinos from  $^{238}\text{U}$  and  $^{232}\text{Th}$  decay chains. The studies of geoneutrinos can help refine the Earth  
646 crust models but is also necessary to characterise their signal, as they are a background to the mass  
647 ordering and oscillations parameters studies.

### 648 Atmospheric neutrinos

649 Atmospheric neutrinos are neutrinos originating from the decay of  $\pi$  and  $K$  particles that are pro-  
650 duced in extensive air showers initiated by the interactions of cosmic rays with the Earth atmosphere.  
651 Earth is mostly transparent to neutrinos below the PeV energy, thus JUNO will be able to see  
652 neutrinos coming from all directions. Their baseline range is large (15km  $\sim$  13000km), they can  
653 have energy between 0.1 GeV and 10 TeV and will contain all neutrino and antineutrinos flavour.  
654 Their studies is complementary to the reactor antineutrinos and can help refine the constraints on  
655 the NMO [60].

656 **Supernovae burst neutrinos**

657 Neutrinos are crucial component during all stages of stellar collapse and explosion. Detection of  
 658 neutrinos coming from core collapse supernovae will provide us important informations on the mech-  
 659 anisms at play in those events. Thanks to its 20 kt sensible volume, JUNO has excellent capabilities  
 660 to detect all flavour of the  $\mathcal{O}(10 \text{ MeV})$  postshock neutrinos, and using neutrinos of the  $\mathcal{O}(1 \text{ MeV})$   
 661 will give informations about the pre-supernovae neutrinos. All those informations will allow to  
 662 disentangle between the multiple hydro-dynamic models that are currently used to describe the  
 663 different stage of core-collapse supernovae.

664 **Diffuse supernovae neutrinos background**

665 Core-collapse supernovae in our galaxy are rare events, but they frequently occur throughout the  
 666 visible Universe sending burst of neutrinos in direction of the Earth. All those events contributes to  
 667 a low background flux of low-energy neutrinos called the Diffuse Supernovae Neutrino Background  
 668 (DSNB). Its flux and spectrum contains informations about the red-shift dependent supernovae rate,  
 669 the average supernovae neutrino energy and the fraction of black-hole formation in core-collapse su-  
 670 pernovae. Depending of the DSNB model, we can expect 2-4 IBD events per year in the energy range  
 671 above the reactor  $\bar{\nu}_e$  signal, which is competitive with the current Super-Kamiokande+Gadolinium  
 672 phase [72].

673 **Beyond standard model neutrinos interactions**

674 JUNO will also be able to probe for beyond standard model neutrinos interactions. After the main  
 675 physics topics have been accomplished, JUNO could be upgraded to probe for neutrinoless beta  
 676 decay ( $0\nu\beta\beta$ ). The detection of such event would give critical informations about the nature of  
 677 neutrinos, is it a majorana or a dirac particle. JUNO will also be able to probe for neutrinos that  
 678 would come for the decay or annihilation of Dark Matter inside the sun and neutrinos from putative  
 679 primordial black hole. Through the unitary test of the mixing matrix, JUNO will be able to search for  
 680 light sterile neutrinos. Thanks to JUNO sensitivity, multiple other exotic research can be performed  
 681 on neutrino related beyond standard model interactions.

682 **Proton decay**

683 Proton decay is a potential unobserved event where the proton decay by violating the baryon num-  
 684 ber. This violation is necessary to explain the baryon asymmetry in the universe and is predicted  
 685 by multiple Grand Unified Theories which unify the strong, weak and electromagnetic interactions.  
 686 Thanks to its large active volume, JUNO will be able to take measurement of the potential proton  
 687 decay channel  $p \rightarrow \bar{\nu}K^+$  [73] thanks to the timing resolution of the SPMT system. Studies show  
 688 that JUNO should be competitive with the current best limit at  $5.9 \times 10^{33}$  years from Super-K. This  
 689 studies show that JUNO, considering no proton decay events observed, would be able to rules a  
 690 limit of  $9.6 \times 10^{33}$  years at 90 % C.L.

691 **2.3 The JUNO detector**

692 The JUNO detector is a scintillator detector buried 693.35 meters under the ground (1800 meters  
 693 water equivalent). It consist of Central Detector (CD), a water pool and a Top Tracker (TT) as showed  
 694 in Figure 2.4a. The CD is an acrylic vessel containing the 20 ktons of Liquid Scintillator (LS). It is  
 695 supported by a stainless steel structure and is immersed in that water pool that is used as shielding

from external radiation and as a cherenkov detector for the background. The top of the experiment is partially covered by the Top Tracker (TT), a plastic scintillator detector which is used to detect the atmospheric muons background and is acting as a veto detector.

The top of the experiment also host the LS purification system, a water purification system, a ventilation system to get rid of the potential radon in the air. The CD is observed by two systems of Photo-Multiplier Tubes (PMT). They are attached to the steel structure and their electronic readout is submersed near them. A third system of PMT is also installed on the structure but are facing outward of the CD, instrumenting the water to be cherenkov detector. The CD and the cherenkov detector are optically separated by Tyvek sheet. A chimney for LS filling and purification and for calibration operations connects the CD to the experimental hall from the top.

The CD has been dimensioned to meet the requirements presented in Section 2.1.1:

- Its 20 ktons monolithic LS provide a volume sizeable enough, in combination with the expected  $\bar{\nu}_e$  flux, to reach the desired statistic in 6 years. Its monolithic nature also allow for a full containment of most of the events, preventing the energy loss in non-instrumented parts that would arise from a segmented detector.
- Its large overburden shield it from most of the atmospheric background that would pollute the signal.
- The localization of the experiment, chosen to maximize the disappearance with a 53km baseline and in a region that allow two nuclear power plant to be used as sources.



(A) Schematics view of the JUNO detector.

(B) Top down view of the JUNO detector under construction

FIGURE 2.4

This section covers in details the different components of the detector and the detection systems.

### 2.3.1 Detection principle

The CD will detect the neutrino and measure their energy mainly via an Inverse Beta Decay (IBD) interaction with proton mainly from the  $^{12}\text{C}$  and H nucleus in the LS:

$$\bar{\nu}_e + p \rightarrow n + e^+$$

Kinematics calculation shows that this interaction has an energy threshold for the  $\bar{\nu}_e$  of  $(m_n + m_e - m_p) \approx 1.806$  MeV [74]. This threshold makes the experiment blind to very low energy neutrinos.

719 The residual energy  $E_\nu - 1.806$  MeV is be distributed as kinetic energy between the positron and the  
 720 neutron. The energy of the emitted positron  $E_e$  is given by [74]

$$E_e = \frac{(E_\nu - \delta)(1 + \epsilon_\nu) + \epsilon_\nu \cos \theta \sqrt{(E_\nu - \delta)^2 + \kappa m_e^2}}{\kappa} \quad (2.2)$$

721 where  $\kappa = (1 + \epsilon_\nu)^2 - \epsilon_\nu^2 \cos^2 \theta \approx 1$ ,  $\epsilon_\nu = \frac{E_\nu}{m_p} \ll 1$  and  $\delta = \frac{m_n^2 - m_p^2 - m_e^2}{2m_p} \ll 1$ . We can see from this  
 722 equation that the positron energy is strongly correlated to the neutrino energy.

723 The positron and the neutron will then propagate in the detection medium, the Liquid Scintillator  
 724 (LS), loosing their kinetic energy by exciting the molecule of the LS (more details in Section 2.3.2).  
 725 Once stopped, the positron will annihilate with an electron from the medium producing two 511  
 726 KeV gamma. Those gamma will themselves interact with the LS, exciting it before being absorbed  
 727 by photoelectrical effect. The neutron will be captured by an hydrogen, emitting a 2.2 MeV gamma  
 728 in the process. This gamma will also deposit its energy before being absorbed by the LS.

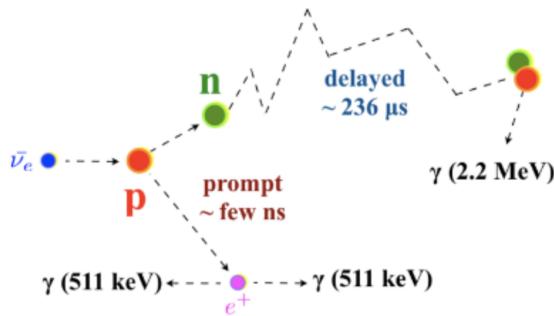


FIGURE 2.5 – Schematics of an IBD interaction in the central detector of JUNO

729 The scintillation photons have frequency in the UV and will propagate in the LS, being re-absorbed  
 730 and re-emitted by compton effect before finally be captured by PMTs instrumenting the acrylic  
 731 sphere. The analog signal of the PMTs digitized by the electronic is the signal of our experiment.  
 732 The signal produced by the positron is subsequently called the prompt signal, and the signal coming  
 733 from the neutron the delayed signal. This naming convention come from the fact that the positron  
 734 will deposit its energy rather quickly (few ns) where the neutron will take a bit more time ( $\sim 236$   $\mu$ s).

### 735 2.3.2 Central Detector (CD)

736 The central detector, composed of 20 ktons of Liquid Scintillator (LS), is the main part of JUNO. The  
 737 LS is contained in a spherical acrylic vessel supported by a stainless steel structure. The CD and  
 738 its structural support are submerged in a cylindrical water pool of 43.5m diameter and 44m height.  
 739 We're confident that the water pool provide sufficient buffer protection in every direction against the  
 740 rock radioactivity.

#### 741 Acrylic vessel

742 The acrylic vessel is a spherical vessel of inner diameter of 35.4 m and a thickness of 120 mm. It is  
 743 assembled from 265 acrylic panels, thermo bonded together. The acrylic recipes has been carefully  
 744 tuned with extensive R&D to ensure it does not include plasticizer and anti-UV material that would  
 745 stop the scintillation photons. Those panels requires to be pure of radioactive materials to not  
 746 cause background. Current setup where the acrylic panels are molded in cleanrooms of class 10000,

let us reach a uranium and thorium contamination of <0.5 ppt. The molding and thermoforming processes is optimized to increase the assemblage transparency in water to >96%. The acrylic vessel is supported by a stainless steel structure via supporting node (fig 2.6). The structure and the nodes are designed to be resilient to natural catastrophic events such as earthquake and can support many times the effective load of the acrylic vessel.

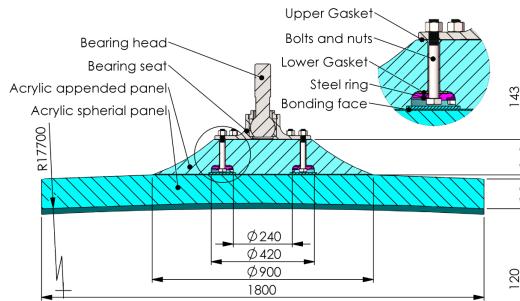


FIGURE 2.6 – Schematics of the supporting node for the acrylic vessel

## 752 Liquid scintillator

The Liquid Scintillator (LS) has a similar recipe as the one used in Daya Bay [75] but without gadolinium doping. It is made of three components, necessary to shift the wavelength of emitted photons to prevent their reabsorption and to shift their wavelength to the PMT sensitivity region as illustrated in Figure 2.7:

- 753 1. The detection medium, the *linear alkylbenzene* (LAB). Selected because of its excellent transparency, high flash point, low chemical reactivity and good light yield. Accounting for ~ 98% of the LS, it is the main component with which ionizing particles and gamma interact. Charged particles will collide with its electronic cloud transferring energy to the molecules, gamma will interact via compton effect with the electronic cloud before finally be absorbed via photoelectric effect.
- 754 2. The second component of the LS is the *2,5-diphenyloxazole* (PPO). A fraction of the excitation energy of the LAB is transferred to the PPO, mainly via non radiative process [76]. The PPO molecules de-excites in the same way, transferring their energy to the bis-MSB. The PPO makes for 1.5 % of the LS.
- 755 3. The last component is the *p-bis(o-methylstyryl)-benzene* (bis-MSB). Once excited by the PPO, it will emit photon with an average wavelength of ~ 430 nm (full spectrum in Figure 2.7) that can thus be detected by our photo-multipliers systems. It amount for ~ 0.5% of the LS.

This formula has been optimized using dedicated studies with a Daya Bay detector [75, 78] to reach the requirements for the JUNO experiment:

- 756 — A light yield / MeV of the amount of  $10^4$  photons to maximize the statistic in the energy measurement.
- 757 — An attenuation length comparable to the size of the detector to prevent losing photons during their propagation in the LS. The final attenuation length is 25.8m [79] to compare with the CD diameter of 35.4m.
- 758 — Uranium/Thorium radiopurity to prevent background signal. The reactor neutrino program require a contamination fraction  $F < 10^{-15}$  while the solar neutrino program require  $F < 10^{-17}$ .

The LS will frequently be purified and tested in the Online Scintillator Internal Radioactivity Investigation System (OSIRIS) [80] to ensure that the requirements are kept during the lifetime of the experiment, more details to be found in Section 2.5.2.

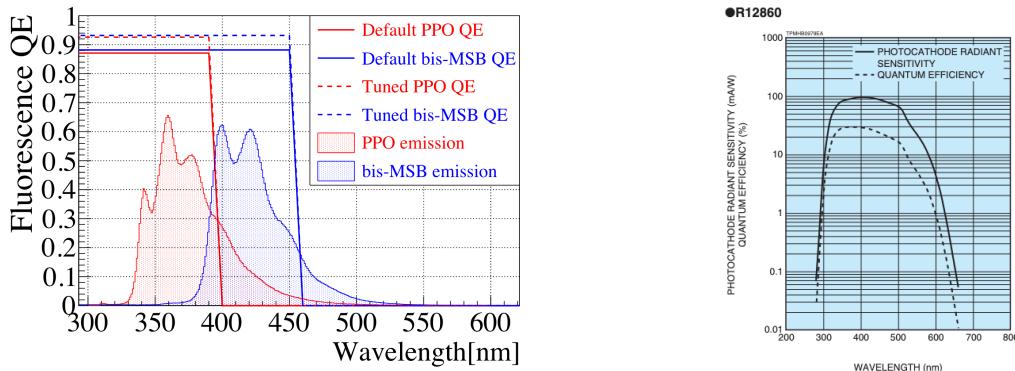


FIGURE 2.7 – On the left: Quantum efficiency (QE) and emission spectrum of the LAB and the bis-MSB [75]. On the right: Sensitivity of the Hamamatsu LPMT depending on the wavelength of the incident photons [77].

### 783 Large Photo-Multipliers Tubes (LPMTs)

784 The scintillation light produced by the LS is then collected by Photo-Multipliers Tubes (PMT) that  
 785 transform the incoming photon into an electric signal. As described in Figure 2.8, the incident  
 786 photons interact with the photocathode via photoelectric effect producing an electron called a Photo-  
 787 Electron (PE). This PE is then focused on the dynodes where the high voltage will allow it to be  
 788 multiplied. After multiple amplification the resulting charge - in coulomb [C] - is collected by the  
 789 anode and the resulting electric signal can be digitalized by the readout electronics from which the  
 790 charge and timing can be extracted.

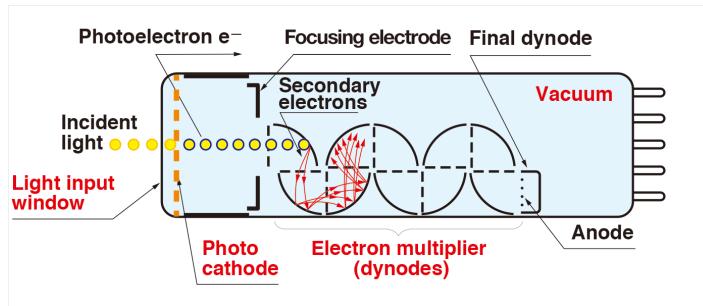


FIGURE 2.8 – Schematic of a PMT

791 The Large Photo-Multipliers Tubes (LPMT), used in the central detector and in the water pool, are  
 792 20-inch (50.8 cm) radius PMTs.  $\sim 5000$  dynode-PMTs [77] were produced by the Hamamatsu<sup>©</sup>  
 793 company and  $\sim 15000$  Micro-Channel Plate (MCP) [81] by the NNVT<sup>©</sup> company. This system is  
 794 the one responsible for the energy measurement with a energy resolution of  $3\%/\sqrt{E}$ , resolution  
 795 necessary for the mass ordering measurement. To reach this precision, the system is composed of  
 796 17612 PMTs quasi uniformly distributed over the detector for a coverage of 75.2% reaching  $\sim 1800$   
 797 PE/MeV or  $\sim 2.3\%$  resolution due to statistic, leaving  $\sim 0.7\%$  for the systematic uncertainties. They  
 798 are located outside the acrylic sphere in the water pool facing the center of the detector. To maintain  
 799 the resolution over the lifetime of the experiment, JUNO require a failure rate  $< 1\%$  over 6 years.

800 The LPMTs electronic are divided in two parts. One "near", located underwater, in proximity of the  
 801 LPMT to reduce the cable length between the PMT and early electronic. A second one, outside of the  
 802 detector that is responsible for higher level analysis before sending the data to the DAQ.

803 The light yield per MeV induce that a LPMT can collect between 1 and 1000 PE per event, a wide

804 dynamic range, causing non linearity in the PMT response that need to be understood and calibrated,  
 805 see Section 2.4 for more details.

806 Before performing analysis, the analog readout of the LPMT need to be amplified, digitised and  
 807 packaged by the readout electronics schematized in Figure 2.9. This electronic is splitted in two  
 808 parts: *wet* electronic that are located near the LPMTs, protected in an Underwater Box (UWB) and  
 809 the *dry* electronics located in deicated rooms outside of the water pool.

810 The LPMTs are connected to the UWB by groups of three. Each UWB contains:

- 811 — Three high voltage units, each one powering a PMT.
- 812 — A global control unit, responsible for the digitization of the waveform, composed of six analog-  
 813 digital units that produce digitized waveform and a Field Programmable Gate Array (FPGA)  
 814 that complete the waveform with metadatas such as the local timestamp trigger, etc... Ths  
 815 FPGA also act as a data buffer when needed by the DAQ and trigger system.
- 816 — Additional memory in order to temporally store the data in case of sudden burst of the input  
 817 rate (such as in the case of nearby supernovae).

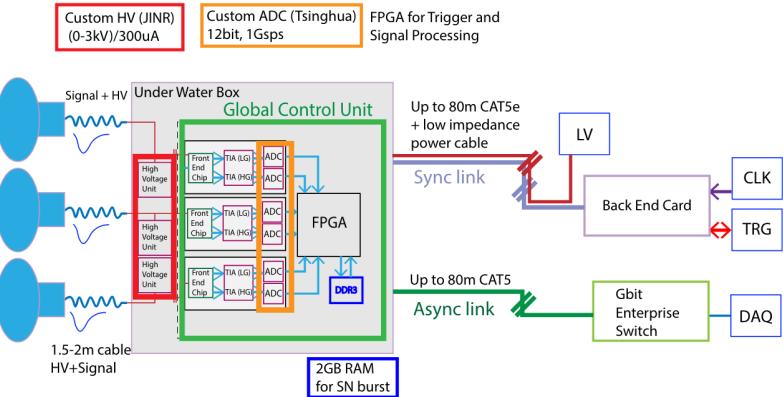


FIGURE 2.9 – The LPMT electronics scheme. It is composed of two part, the *wet* electronics on the left, located underwater and the *dry* electronics on the right. They are connected by Ethernet cable for data transmission and a dedicated low impedance cable for power distribution

818 The *dry* electronic synchronize the signals from the UWBs abd centralise the information of the CD  
 819 LPMTs. It act as the Global Trigger by sending the UWB data to DAQ in the case if the LPMT  
 820 multiplicity condition is fulfilled.

### 821 Small Photo-Multipliers Tubes (SPMTs)

822 The Small PMT (SPMTs) system is made of 3-inch (7.62 cm) PMTs. They will be used in the CD  
 823 as a secondary detection system. Those 25600 SPMTs will observe the same events as the LPMTs,  
 824 thus sharing the physics and detector systematics up until the photon conversion. With a detector  
 825 coverage of 2.7%, this system will collect  $\sim 43$  PE/MeV for a final energy resolution of  $\sim 17\%$ .  
 826 This resolution is not enough to measure the NMO,  $\theta_{13}$ ,  $\Delta m_{31}^2$  but will be sufficient to independently  
 827 measure  $\theta_{12}$  and  $\Delta m_{21}^2$ .

828 The benefit of this second system is to be able to perform another, independent measure of the  
 829 same events as the LPMTs, constituting the Dual Calorimetry useful for calibrationa and, as it we  
 830 will explore in this thesis, for physics analysis. Due to the low PE rate, SPMTs will be running in  
 831 photo-counting mode in the reactor range and thus will be insensitive to LPMT intrinsic effect (see

Section 2.4). Using this property, the intrinsic charge non linearity of the LPMTs can be measured by comparing the PE count in the SPMTs and LPMTs [82]. Also, due to their smaller size and electronics, SPMTs have a better timing resolutions than the LPMTs. At higher energy range, like supernovae events, LPMTs will saturate where SPMTs due to their lower PE collection will to produce a reliable measure of the energy spectrum.

The SPMTs will be grouped by pack of 128 to an UWB hosting their electronics as illustrated in Figure 2.10. This underwater box host two high voltage splitter boards, each one supplying 64 SPMTs, an ASIC Battery Card (ABC) and a global control unit.

The ABC board will readout and digitize the charge and time of the 128 SPMTs signals and a FPGA will joint the different metadata. The global control unit will handle the powering and control of the board and will be in charge of the transmission of the data to the DAQ.

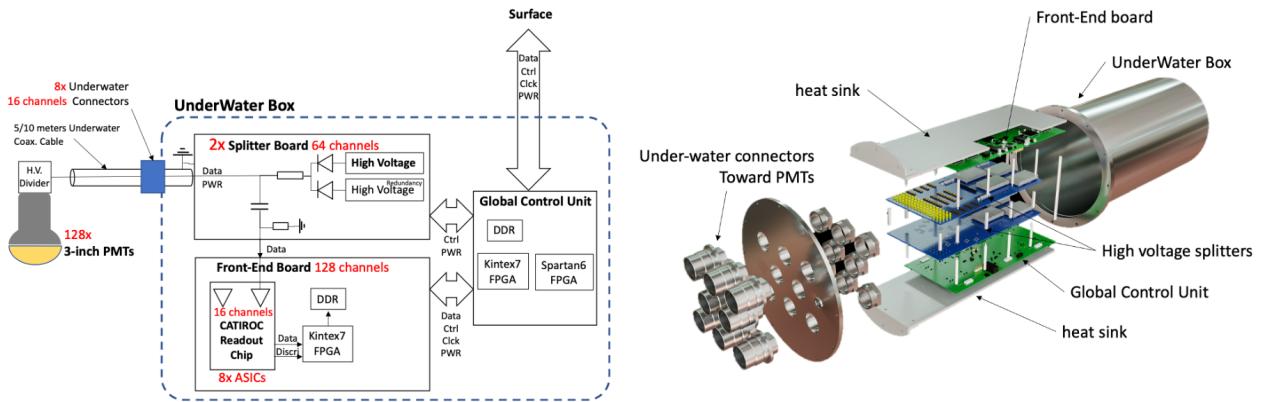


FIGURE 2.10 – Schematic of the JUNO SPMT electronic system (left), and exploded view of the main component of the UWB (right)

### 2.3.3 Veto detector

The CD will be bathed in constant background noise coming from numerous sources : the radioactivity from surrounding rock and its own components or from the flux of cosmic muons. This background needs to be rejected to ensure the purity of the IBD spectrum. To prevent a big part of them, JUNO use two veto detector that will tag events as background before CD analysis.

#### Cherenkov in water pool

The Water Cherenkov Detector (WCD) is the instrumentation of the water buffer around the CD. When high speed charged particles will pass through the water, they will produced cherenkov photons. The light will be collected by 2400 MCP LPMTs installed on the outer surface of the CD structure. The muons veto strategy is based on a PMT multiplicity condition. WCD PMTs are grouped in ten zones: 5 in the top, 5 in the bottom. A veto is raised either when more than 19 PMTs are triggered in one zone or when two adjacent zones simultaneously trigger more than 13 PMTs. Using this trigger, we expect to reach a muon detection efficiency of 99.5% while keeping the noise at reasonable level.

857 **Top tracker**

858 The JUNO Top Tracker (TT) is a plastic scintillator detector located on the top of the experiment (see  
 859 Figure 2.11). Made from plastic scintillator from OPERA [83] layered horizontally in 3 layers on the  
 860 top of the detector, the TT will be able to detect incoming atmospheric muons. With its coverage,  
 861 about 1/3 of the of all atmospheric muons that passing through the CD will also pass through the 3  
 862 layer of the detector. While it does not cover the majority of the CD, the TT is particularly effective to  
 863 detect muons coming through the filling chimney region which might present difficulties from the  
 other subsystems in some classes of events.

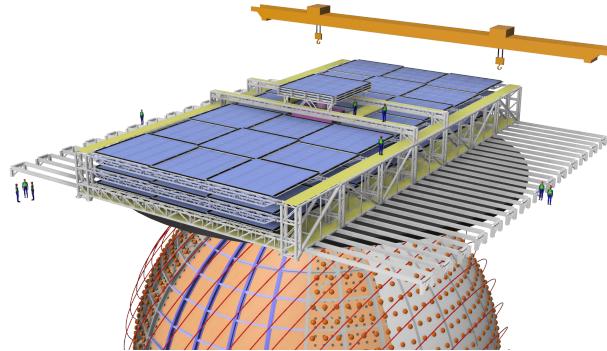


FIGURE 2.11 – The JUNO top tracker

864

865 **2.4 Calibration strategy**

866 The calibration is a crucial part of the JUNO experiment. The detector will continuously bath in  
 867 neutrinos coming from the close nuclear power plant, from other sources such as geo neutrinos,  
 868 the sun and will be exposed to background noise coming from atmospheric muons and natural  
 869 radioactivity. Because of this continuous rate, low frequency signal event, we need high frequency,  
 870 recognisable sources in the energy range of interest : [0-12] MeV for the positron signal and 2.2 MeV  
 871 for the neutron capture. It is expected that the CD response will be different depending on the type  
 872 of particle, due to the interaction with LS, the position on the event and the optical response of the  
 873 acrylic sphere (see Section 3.3). We also expect a non-linear energy response of the CD due to the LS  
 874 properties [75] but also due to the reponse of the LPMTs system when collecting a large amount of  
 875 PE [82].

876 **2.4.1 Energy scale calibration**

877 While electrons and positrons sources would be ideal, for a large LS detector thin-walled electrons  
 878 or positrons sources could lead to leakage of radionucleides causing radioactive contamination.  
 879 Instead, we consider gamma sources in the range of the prompt energy of IBDs. The sources are  
 880 reported in table 2.3.

881 For the  $^{68}\text{Ge}$  source, it will decay in  $^{68}\text{Ga}$  via electron capture, which will itself  $\beta^+$  decay into  $^{68}\text{Zn}$ .  
 882 The positrons will be absorbed by the enclosure so only the annihilation gamma will be released. In  
 883 addition,  $(\alpha, n)$  sources like  $^{241}\text{Am-Be}$  and  $^{241}\text{Am-}^{13}\text{C}$  are used to provide both high energy gamma  
 884 and neutrons, which will later be captured in the LS producing the 2.2 MeV gamma.

Sources / Processes	Type	Radiation
$^{137}\text{Cs}$	$\gamma$	0.0662 MeV
$^{54}\text{Mn}$	$\gamma$	0.835 MeV
$^{60}\text{Co}$	$\gamma$	$1.173 + 1.333$ MeV
$^{40}\text{K}$	$\gamma$	1.461 MeV
$^{68}\text{Ge}$	$e^+$	annihilation 0.511 + 0.511 MeV
$^{241}\text{Am-Be}$	$n, \gamma$	neutron + 4.43 MeV ( $^{12}\text{C}^*$ )
$^{241}\text{Am-}^{13}\text{C}$	$n, \gamma$	neutron + 6.13 MeV ( $^{16}\text{O}^*$ )
$(n, \gamma)p$	$\gamma$	2.22 MeV
$(n, \gamma)^{12}\text{C}$	$\gamma$	4.94 MeV or 3.68 + 1.26 MeV

TABLE 2.3 – List of sources and their process considered for the energy scale calibration

885 From this calibration we call  $E_{\text{vis}}$  the "visible energy" that is reconstructed by our current algorithms  
 886 and we compare it to the true energy deposited by the calibration source. The results shown in  
 887 Figure 2.12 show the expected response of the detector from calibration sources. The non-linearity is  
 888 clearly visible from the  $E_{\text{vis}}/E_{\text{true}}$  shape. See [84] for more details.

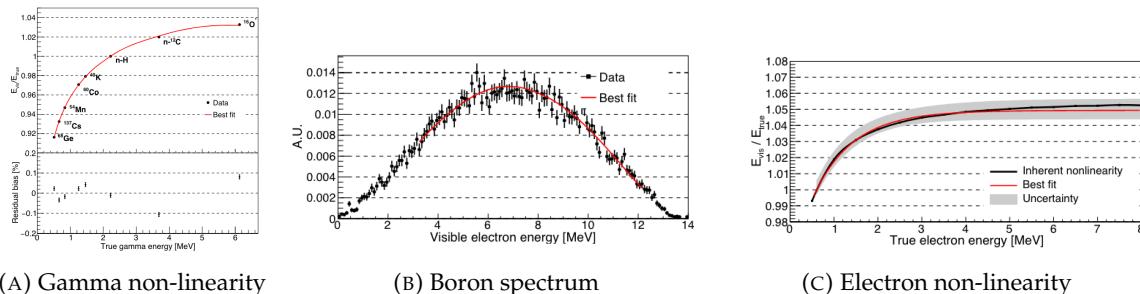


FIGURE 2.12 – Fitted and simulated non linearity of gamma, electron sources and from the  $^{12}\text{B}$  spectrum. Black points are simulated data. Red curves are the best fits. Figures taken from [84].

## 2.4.2 Calibration system

890 The non-uniformity due to the event position in the detector (more details in Section 3.3) will be  
 891 studied using multiples systems that are schematized in Figure 2.13. They allow to position sources  
 892 at different location in the CD.

- 893 — For a one-dimension vertical calibration, the Automatic Calibration Unit (ACU) will be able to  
 894 deploy multiple radioactive sources or a pulse laser diffuser ball along the central axis of the  
 895 CD through the top chimney. The source position precision is less than 1cm.
- 896 — For off-axis calibration, a calibration source attached to a Cable Loop System (CLS) can be  
 897 moved on a vertical half-plane by adjusting the length of two connection cable. Two set of CSL  
 898 will be deployed to provide a 79% effective coverage of a vertical plane.
- 899 — A Guiding Tube (GT) will surround the CD to calibrate the non-uniformity of the response at  
 900 the edge of the detector
- 901 — A Remotely Operated under-LS Vehicle (ROV) can be deployed to desired location inside LS  
 902 for a more precise and comprehensive calibration. The ROV will also be equipped with a  
 903 camera for inspection of the CD.

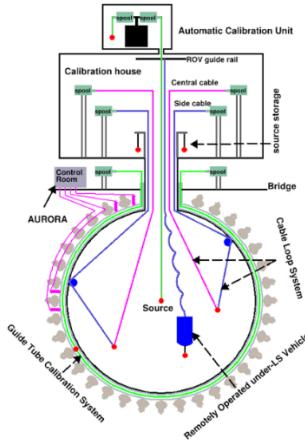


FIGURE 2.13 – Overview of the calibration system

904 The preliminary calibration program is depicted in table 2.4.

Program	Purpose	System	Duration [min]
Weekly calibration	Neutron (Am-C)	ACU	63
	Laser	ACU	78
Monthly calibration	Neutron (Am-C)	ACU	120
	Laser	ACU	147
	Neutron (Am-C)	CLS	333
	Neutron (Am-C)	GT	73
Comprehensive calibration	Neutron (Am-C)	ACU, CLS and GT	1942
	Neutron (Am-Be)	ACU	75
	Laser	ACU	391
	$^{68}\text{Ge}$	ACU	75
	$^{137}\text{Cs}$	ACU	75
	$^{54}\text{Mn}$	ACU	75
	$^{60}\text{Co}$	ACU	75
	$^{40}\text{K}$	ACU	158

TABLE 2.4 – Calibration program of the JUNO experiment

### 905 2.4.3 Instrumental non-linearity calibration

906 One of the main interests of Dual Calorimetry is to calibrate away an instrumental effect called charge  
907 non linearity (QNL), which will be described in more detail in Chapter 7.

908 In short, during a typical IBD event, between 0 and 100 PEs can be produced in a given LPMT  
909 (depending on the position of the interaction and the positron energy). This is a large dynamic range.  
910 When the number of PEs is high, the reconstruction of the LPMT charge can become inaccurate,  
911 underestimating the actual number of PEs as illustrated in Figure 2.14. This QNL is difficult to  
912 separate from other non linearities (like the non linearity in the LS photon yield as a function of the  
913 deposit energy). In chapter 5 and 6 of this thesis [82], a calibration method that constitutes the core of  
914 dual calorimetry are described. They are based on the comparisons between signals seen in LPMTs  
915 and signals seen in SPMTs. In the latter system, due to its small angular coverage, individual SPMT  
916 rarely see more than 1 PE per event, and therefore are essentially immune against QNL. The method  
917 described in [82] uses a tunable light source covering the range of 0 to 100 PE per LPMT channel

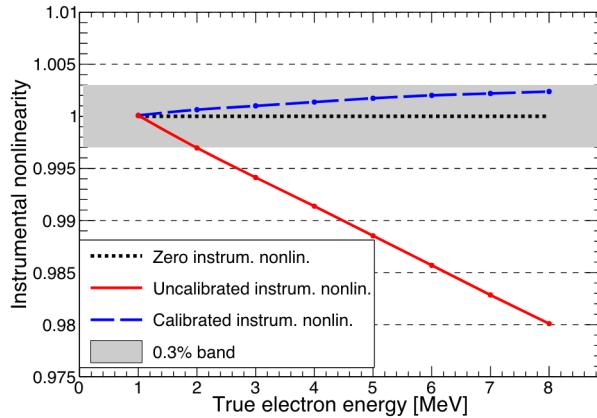


FIGURE 2.14 – Event-level instrumental non-linearity, defined as the ratio of the total measured LPMT charge to the true charge for events at the center of the detector. The solid red line represents event-level non-linearity without the channel-level correction in an extreme hypothetical scenario of 50% non-linearity over 100 PEs for the LPMTs. The dashed blue line represents that after the channel-level correction. The gray band shows the residual uncertainty of 0.3%, after the channel-level correction. Figure taken from [84].

## 918 2.5 Satellite detectors

919 As introduced in Section 2.1.1 and section 2.3.2, the precise knowledge and understanding of the  
 920 detector condition is crucial for the measurements of the NMO and oscillation parameters. Thus  
 921 two satellite detectors will be setup to monitor the experiment condition. TAO to monitor and  
 922 understand the  $\bar{\nu}_e$  flux and spectrum coming from the nuclear reactor and OSIRIS to monitor the  
 923 LS response.

### 924 2.5.1 TAO

925 The Taishan Antineutrino Observatory (TAO) [71, 85] is a ton-level gadolinium doped liquid scin-  
 926 tillator detector that will be located near the Taishan-1 reactor. It aim to measure the  $\bar{\nu}_e$  spectrum at  
 927 very low distance (44m) from the reactor to measure a quasi-unosculated spectrum. TAO also aim  
 928 to provide a major contribution to the so-called reactor anomaly [70]. Its requirement are to the level  
 929 of 2 % energy resolution at 1 MeV.

#### 930 Detector

931 The TAO detector is close, in concept, to the CD of JUNO. It is composed of an acrylic vessel  
 932 containing 2.8 tons of gadolinium-loaded LS instrumented by an array of silicon photomultipliers  
 933 (SiPM) reaching a 95% coverage. To efficiently reduce the dark count of those sensors, the detector is  
 934 cooled to -50 °C. The  $\bar{\nu}_e$  will interact with the LS via IBD, producing scintillation light, that will  
 935 be detected by the SiPMs. From this signal the  $\bar{\nu}_e$  energy and the full spectrum reconstructed.  
 936 This spectrum will then be used by JUNO to calibrate the unoscillated spectrum, most notably the  
 937 fission product fraction that impact the rate and shape of the spectrum. A schema of the detector is  
 938 presented in Figure 2.15a.

### 2.5.2 OSIRIS

The Online Scintillator Internal Radioactivity Investigation System (OSIRIS) [80] is an ultralow background,  $20\text{ m}^3$  LS detector that will be located in JUNO cavern. It aim to monitor the radioactive contamination, purity and overall response of the LS before it is injected in JUNO. OSIRIS will be located at the end of the purification chain of JUNO, monitoring that the purified LS meet the JUNO requirements. The setup is optimized to detect the fast coincidences decay of  $^{214}\text{Bi} - ^{214}\text{Po}$  and  $^{212}\text{Bi} - ^{212}\text{Po}$ , indicators of the decay chains of U and Th respectively.

#### Detector

OSIRIS is composed of an acrylic vessel that will contains 17t of LS. The LS is instrumented by a PMT array of 64 20 inch PMTs on the top and the side of the vessel. To reach the necessary background level required by the LS purity measurements, in addition to being 700m underground in the experiment cavern, the acrylic vessel is immersed in a tank of ultra pure water. The water is itself instrumented by another array of 20 inch PMTs, acting as muon veto. A schema of the detector is presented in Figure 2.15b.

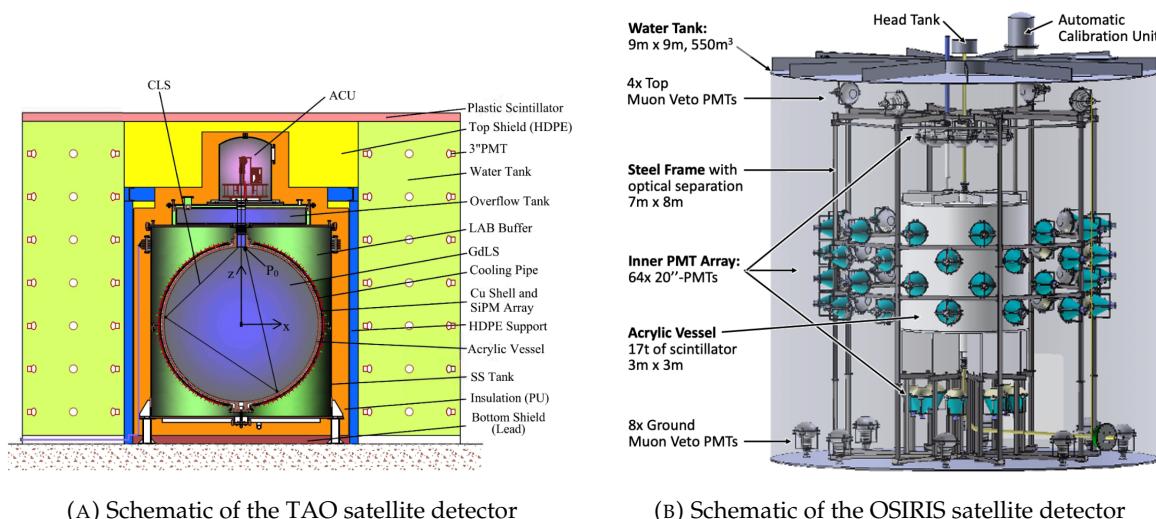


FIGURE 2.15

## 2.6 Software

The simulation, reconstruction and analysis algorithms are all packaged in the JUNO software, subsequently called the software. It is composed of multiple components integrated in the SNiPER [86] framework:

- Various primary particles simulators for the different kind of events, background and calibration sources.
- A Geant4 [87–89] Monte Carlo (MC) simulation containing the detectors geometries, a custom optical model for the LS and the supporting structures of the detectors. The Geant4 simulation integrate all relevant physics process for JUNO, validated by the collaboration. This step of the simulation is commonly called *Detsim* and compute up to the production of photo-electrons in

the PMTs. The optics properties of the different materials and detector components have been measured beforehand to be used to define the material and surfaces in the simulation.

- An electronic simulation, simulating the response waveform of the PMTs, tracking it through the digitization process, accounting for effects such as non-linearity, dark noise, Time Transit Spread (TTS), pre-pulsing, after-pulsing and ringing of the waveform. It's also the step handling the event triggers and mixing. This step is commonly referenced as *ElecSim*.
- A waveform reconstruction where the digitized waveform are filtered to remove high-frequency white noise and then deconvoluted to yield time and charge informations of the photons hits on the PMTs. This step is commonly referenced as *Calib*.
- The charge and time informations are used by reconstruction algorithms to reconstruct the interaction vertex and the deposited energy. This step is commonly reported as *Reco*. See Section 3.3 for more details on the reconstruction.
- Once the singular events are reconstructed, they go through event pairing and classification to select IBD events. This step is named Event Classification.
- The purified signal is then analysed by the analysis framework which depend of the physics topic of interest. An introduction to the reactor  $\bar{\nu}n_e$  is presented in Section 2.7.

The steps Reco and Event Classification are divided into two category of algorithm. Fast but less accurate algorithms that are running during the data taking designated as the *Online* algorithms. Those algorithm are used to take the decision to save the event on tape or to throw it away. More accurate algorithms that run on batch of events designated *Offline* algorithms. They are used for the physics analysis. The Offline Reco will be one of the main topic of interest for this thesis.

## 2.7 Reactor anti-neutrino oscillation analysis

### 2.7.1 IBD samples selection

The  $\bar{\nu}_e$  coming from nuclear reactor will, for the most part, interact with proton, hydrogen nucleus, via Inverse Beta Decay (IBD). The first step of the oscillation analysis is to constitute a sample of IBD candidates, dominated by actual IBDs. The IBD interaction, schematised in Figure 2.5, will produce two particle, with differentiable signals.

The first signal comes from the positron slowdown and its annihilation with an electron of the LS. This is the *prompt* signal, happening a few ns after the IBD. The positron takes most of the  $\bar{\nu}_e$  kinetic energy, as detailed in Section 2.3.1.

The leftover kinetic energy is taken by the neutron that, after thermalisation in the LS, will be captured by an hydrogen and produce a 2.2 MeV gamma, or by a carbon emitting a 4.9 MeV gamma. This is the *delayed* signal, happening  $\sim$ 236  $\mu$ s after the IBD. This second mono-energetic event serve as a marker for the IBD.

The IBD selection is thus based on the selection of a prompt event, with an energy between 0.8 and 12 MeV, and a delayed event with an energy in the ranges [1.9, 2.5] MeV or [4.4, 5.5] MeV. Those two signal needs to be in a 1 ms time window and within 1.5 m from each other. Additionally the two signal needs to be in a radius of 17.2m from the detector center (0.5 m from the edge) to protect from accidental background formed by two uncorrelated signals [90]. Those values will be further refined after once JUNO data-taking starts.

In addition, specials veto are setup to protect from cosmic muons and their aftermath. The details of those veto and selection can be found in [90].

1005 The expected rate and selection efficiency on IBD can be found in table 2.5. After these selection, the  
 1006 residual background, including  $\bar{\nu}_e$  coming from other sources than the reactor can be found in table  
 1007 2.6.

Selection Criterion	Efficiency [%]	IBD Rate [day <sup>-1</sup> ]
All IBDs	100.0	57.4
Fiducial Volume	91.5	52.5
IBD Selection	98.1	51.5
Energy Range	99.8	-
Time Correlation ( $\Delta T_{p-d}$ )	99.0	-
Spatial Correlation ( $\Delta R_{p-d}$ )	99.2	-
Muon Veto (Temporal + Spatial)	91.6	47.1
Combined Selection	82.2	47.1

TABLE 2.5 – Summary of cumulative reactor antineutrino selection efficiencies. The reported IBD rates (with baselines <300 km) refer to the expected events per day after the selection criteria are progressively applied. Table taken from [90]

Backgrounds	Rate [day <sup>-1</sup> ]	B/S [%]
Geoneutrinos	1.2	2.5
World reactors	1.0	2.1
Accidentals	0.8	1.7
<sup>9</sup> Li/ <sup>8</sup> He	0.8	1.7
Atmospheric neutrinos	0.16	0.34
Fast neutrons	0.1	0.21
<sup>13</sup> C( $\alpha, n$ ) <sup>16</sup> O	0.05	0.01
Total backgrounds	4.11	8.7

TABLE 2.6 – Expected background rates, background to signal ratio (B/S), and rate and shape uncertainties. The B/S ratio is calculated by using the IBD signal rate of 47.1/day. Table taken from [90]

1008 Once a sample is obtained, the oscillation analysis will consist essentially on the fit of a spectrum  
 1009 model to the spectrum observed in the selected sample. More specifically, the spectrum under  
 1010 analysis is the spectrum of the reconstructed visible energy of the positron :  $E_{vis}^{e^+}$ . The reconstruction  
 1011 is presented in detail in Section 3.3. For 6 years of data taking, it will resemble that on Figure 2.3.  
 1012 In the next sections, I describe the fit procedures developed in JUNO. This will be the occasion to  
 1013 introduce notions useful for Chapter 7. Besides, I'll also describe the versions of the fit used in this  
 1014 Chapter 7.

### 1015 2.7.2 Synthetic overview of fit procedures developed at JUNO

1016 Several fit procedures are being developed by JUNO collaborators (half a dozen of groups work  
 1017 in parallel within the collaboration). We do not have the ambition of a thorough description here.  
 1018 Instead, we try to introduce the main elements useful to the reader to understand JUNO's future  
 1019 results, and the fit procedures used Chapter 7.

1020 In most cases, the fit is a binned fit to the histogrammed spectrum of  $E_{vis}^{e^+}$ , like the one in Figure 2.3.  
 1021 It is based on the minimization of a  $\chi^2$  test statistics. Generically, it can be written this way :

$$\chi^2 = (\mathbf{T}(\boldsymbol{\theta}, \boldsymbol{\eta}) - \mathbf{D})^T \mathbf{V}^{-1} (\mathbf{T}(\boldsymbol{\theta}, \boldsymbol{\eta}) - \mathbf{D}) + \chi^2_{nuis}(\boldsymbol{\eta}) \quad (2.3)$$

1022 where the components of data vector  $\mathbf{D}$  are the number of events found in individual bins of the

fitted histogram,  $T(\boldsymbol{\theta}, \boldsymbol{\eta})$  is the vector of the predicted number of entries in each bins. This prediction is the integration over the width of the bins of the spectrum model for a given NMO (described latter in this section).

This model depends on the oscillation parameters  $\boldsymbol{\theta} = (\Delta m_{21}^2, \sin^2(2\theta_{12}), \Delta m_{31}^2, \sin^2(2\theta_{13}))$ , and on nuisance parameters  $\boldsymbol{\eta}$  involved in the fit model and associated with systematic uncertainties. Uncertainties are treated in two ways : statistical and some of the systematic uncertainties are accounted for via the covariance matrix  $V = V_{stat} + V_{syst}$ ; remaining systematic uncertainties are treated via the penalty term  $\chi^2_{nuis}$ , which is written this way :

$$\chi^2_{nuis}(\boldsymbol{\eta}) = (\boldsymbol{\eta} - \bar{\boldsymbol{\eta}})^T \cdot V_{\boldsymbol{\eta}}^{-1}(\boldsymbol{\eta}) \cdot (\boldsymbol{\eta} - \bar{\boldsymbol{\eta}}) \quad (2.4)$$

where  $\bar{\boldsymbol{\eta}}$  is the vector containing the most probable values of the nuisance parameters according to our knowledge prior to the fit, and where  $V_{\boldsymbol{\eta}}$  is the covariance matrix accounting of the uncertainty on these values, and the potential correlations between them. In principles, a likelihood could be used instead of a  $\chi^2$ . However, some of the systematic uncertainties are not trivial to parameterize, therefore treating them as nuisance parameters in not trivial.

An example of nuisance parameters are the  $A$ ,  $B$  and  $C$  parameters of equation 7.19, which can be used to describe the resolution on the reconstructed energy. The fit model leading to  $T(\boldsymbol{\theta}, \boldsymbol{\eta})$  indeed incorporates this resolution.

### 1039 Treatment of uncertainties

Differences between various fit procedures developed within JUNO often lies in the choice of the systematic uncertainties that are treated via  $V$  or  $\chi^2_{nuis}(\boldsymbol{\eta})$ . Among the reasons behind these differences is the necessity to compare several approaches to ensure the robustness JUNO's oscillation analysis results. This approach was already adopted in the recent evaluations of JUNO's potential [61, 90]. Studies carried out so far at Subatech assumes a treatment entirely via  $V$ .

Other differences lies in the choice of the way to evaluate  $V_{stat}$ . Two common approaches used in  $\chi^2$  fit are the Neyman and the Pearson approaches. If the size of the fitted sample is high enough, the variation of  $D_i$ , the number of entries in bin  $i$ , around its true expectation value  $\bar{D}_i$  is  $\sqrt{\bar{D}_i}$ . To evaluate this number, the Neyman approach uses simply the number of entries observed in the sample under analysis :  $\sqrt{D_i}$ . The Pearson approach uses the prediction by the fit model :  $\sqrt{T(\boldsymbol{\theta}, \boldsymbol{\eta})_i}$ .

Both cases are approximations which lead to biases that are not tolerable given the precision JUNO must aim at for a successful oscillation analysis. To reduce this bias, most of JUNO groups employ the "Combined Neyman Pearson" approach introduced in [91]. Schematically, it consists on combining both approaches :  $(V_{stat})_{ii} = 3 / \left( \frac{1}{T(\boldsymbol{\theta}, \boldsymbol{\eta})_i} + \frac{2}{D_i} \right)$ . Weights in this relation are chosen in order to cancel typical biases. The validity of this method is not guaranteed universally. In particular, limitations appear when a complex systematic matrix  $V_{syst}$  is added to  $V_{stat}$ .

This is the case in the approach followed at Subatech, were all sources of systematic uncertainties are treated via this matrix. Dedicated studies run at Subatech observed biases in the fitted oscillation parameters using CNP in this case. Subatech's group therefore adopted another approach (verified to be unbiased).

Originally, fitting the  $E_{vis}^{e+}$  spectrum should mean maximising a likelihood, equal to the product over all bins of the probabilities to find  $D_i$  in bin  $i$ . With a large enough samples, this product tends to a multidimensional gaussian (one dimension per bin) :

$$\mathcal{L} = 2\pi^{-\frac{N}{2}} |V|^{-\frac{1}{2}} e^{-\frac{1}{2}(\mathbf{D} - \mathbf{T}(\boldsymbol{\theta}, \boldsymbol{\eta}))^T V^{-1}(\mathbf{D} - \mathbf{T}(\boldsymbol{\theta}, \bar{\boldsymbol{\eta}}))} \quad (2.5)$$

<sup>1063</sup> Replacing  $\mathcal{L}$  by  $-2 \ln \mathcal{L}$  one obtains :

$$\chi^2_{PV} = (\mathbf{T}(\boldsymbol{\theta}, \boldsymbol{\eta}) - \mathbf{D})^T V^{-1} (\mathbf{T}(\boldsymbol{\theta}, \boldsymbol{\eta}) - \mathbf{D}) + \ln(|V|) \quad (2.6)$$

<sup>1064</sup> where  $V$  is the total covariance matrix with its statistical component evaluated according to the  
<sup>1065</sup> Pearson approach. The  $\ln |V|$  term, often neglected in  $\chi^2$  fits, ensures that biases, essentially related  
<sup>1066</sup> to the normalisation of the fitted distribution, are avoided. This "PearsonV"  $\chi^2$  is the one that we  
<sup>1067</sup> minimize in the fits used in Chapter 7.

<sup>1068</sup> Another difference between the various procedures developed at JUNO is the choice of the spectrum  
<sup>1069</sup> range and binning. So far, at Subatech, we use an histogram defined between 0.8 and 9 MeV, and a  
<sup>1070</sup> regular binning involving 20 keV wide bins.

### <sup>1071</sup> Joint fit of JUNO and TAO spectra

<sup>1072</sup> Another difference between the various fit procedures developed in the collaboration is the inclusion  
<sup>1073</sup> of the data collected by TAO (see Section 2.5.1). The spectrum prediction  $\mathbf{T}(\boldsymbol{\theta}, \boldsymbol{\eta})$  involves predictions  
<sup>1074</sup> on the differential flux of  $\bar{\nu}_e$  as a function of  $E_{\bar{\nu}_e}$  produced in reactors. This is one of the main  
<sup>1075</sup> systematic uncertainties affecting the oscillation analysis. This can be constrained using the data  
<sup>1076</sup> of TAO. An efficient way to use them is via a simultaneous fit, which will constrain the part of the  $\boldsymbol{\eta}$   
<sup>1077</sup> parameters related to the reactor predictions. In this case, equation 2.3 becomes :

$$\chi^2 = \sum_d \left( \mathbf{T}^d(\boldsymbol{\theta}^d, \boldsymbol{\eta}) - \mathbf{D}^d \right)^T V^{-1} \left( \mathbf{T}^d(\boldsymbol{\theta}^d, \boldsymbol{\eta}) - \mathbf{D}^d \right) + \chi^2_{nuis}(\boldsymbol{\eta}) \quad (2.7)$$

<sup>1078</sup> where the  $d$  superscript stands for the spectrum measured in JUNO or TAO.

<sup>1079</sup> Finally, it must be noted that JUNO's sensitivity to  $\sin^2(2\theta_{13})$  is too weak for a competitive measure-  
<sup>1080</sup> ment. In most versions of the oscillation analyses carried out within JUNO, it will be considered as a  
<sup>1081</sup> nuisance parameter. In practice, the various  $\chi^2$ 's presented earlier will receive an additional term :

$$\chi^2_{\sin^2(2\theta_{13})} = \frac{(\sin^2(2\theta_{13}) - \overline{\sin^2(2\theta_{13})})^2}{\sigma_{\sin^2(2\theta_{13})}^2} \quad (2.8)$$

<sup>1082</sup> where  $\overline{\sin^2(2\theta_{13})}$  and the denominators can be provided, for instance, by the world average on this  
<sup>1083</sup> parameter.

### <sup>1084</sup> 2.7.3 The spectrum model and sources of systematic uncertainties

<sup>1085</sup> The  $E_{vis}^{e+}$  spectrum observed in data (Fig 2.3) is the sum of the IBD spectrum and of the various  
<sup>1086</sup> backgrounds spectra (see table 2.6). The spectrum prediction  $\mathbf{T}(\boldsymbol{\theta}, \boldsymbol{\eta})$  is therefore the sum of IBD and  
<sup>1087</sup> backgrounds predictions. The latter are provided by MC simulations. The former results from the  
<sup>1088</sup> theoretical description of the series of phenomena that lead to the observed IBD spectrum. In a given  
<sup>1089</sup> bin  $i$ , it can be expressed this way :

$$T^i(\boldsymbol{\theta}, \boldsymbol{\eta}) = \sum_j C_{ij}^{E_{rec}} \int_{E_j^{vis}}^{E_{j+1}^{vis}} dE^{vis} \int_{-1}^1 d\cos\theta \Phi(E^\nu) \frac{d\sigma}{d\cos\theta}(E^\nu, \cos\theta) \frac{dE^\nu}{dE^{dep}} \frac{dE^{dep}}{dE^{vis}} \quad (2.9)$$

<sup>1090</sup> In the above equation, 4 kinds of energies appears: following the IBD, the antineutrino energy  $E^\nu$  is  
<sup>1091</sup> quasi entirely transferred to the positron, of energy  $E_e$ . It eventually annihilates, so the actual energy  
<sup>1092</sup> released in the LS is  $E_{dep}$ , which includes the mass of the annihilated electron. The production optical

1093 photons is not linear in  $E_{dep}$  (see Section 2.4), so that the visible energy (that will be reconstructed) is  
 1094  $E_{vis}$ . This reconstruction comes with resolution effects, leading to  $E_{rec}$ .

1095 Equation 2.9 describe the passage from the original differential flux (as a function of  $E^\nu$ ) of antineu-  
 1096 trinos reaching the detector to the reconstructed spectrum:

- 1097 —  $\Phi(E^\nu)$  is the differential antineutrino flux reaching JUNO.
- 1098 —  $\frac{d\sigma}{d\cos\theta}(E^\nu, \cos\theta)$  account for the IBD cross section, which depends on the antineutrino energy  
 1099 and on the incidence angle.
- 1100 — The last two terms of the integrand are the differential relations linking  $E^\nu$ ,  $E^{dep}$  and  $E^{vis}$ .
- 1101 — Reconstruction effects are described via  $C_{ij}^{rec}$ 's, that make the link between the true and recon-  
 1102 structed visible energy. In a simple case, it is equivalent to a convolution product. The matrix  
 1103 formalism here prepares the fact that a realistic analysis might employ a more empirical way,  
 1104 based on MC.

1105

1106 The differential flux is expressed this way:

$$\Phi(E^\nu) = \sum_r \left( \frac{\mathcal{P}_{\bar{\nu}_e \rightarrow \bar{\nu}_e}(E^\nu, L_r)}{4\pi L_r^2} \frac{W_r}{\sum_i f_{i,r} e_i} \sum_i f_{i,r} s_i(E^\nu) \right) \quad (2.10)$$

1107 where:

- 1108 —  $\mathcal{P}_{\bar{\nu}_e \rightarrow \bar{\nu}_e}(E^\nu, L_r)$  is the antineutrino survival probability at distance  $L_r$  from the production point  
 1109 in reactor  $r$ , dictated by the oscillation probability.
- 1110 —  $e_i$  stands for the mean energy released per fission for isotope  $i$ .
- 1111 —  $W_r$  is the thermal power of reactor  $r$ .
- 1112 —  $f_{i,r}$  is the fission fraction in reactor  $r$  of isotope  $i$  among the four.
- 1113 —  $s_i(E^\nu)$  is the  $\bar{\nu}_e$  energy spectrum - at emission point - per fission for each isotope, as emitted by  
 1114 the reactor.

1115

## 1116 Sources of systematic uncertainties

1117 The numerous quantities appearing in the spectrum model embody a good part of the systematic  
 1118 uncertainties. Among the leading contributions are those related to the knowledge of the reactor  
 1119 related quantities. Of importance are also the uncertainties related to the modelling of the non  
 1120 linearity of the photon emission (passage from  $E^{dep}$  to  $E^{vis}$ ) and of the reconstruction resolution.  
 1121 The shape and rate of the backgrounds are also a leading source of systematic uncertainties. The  
 1122 uncertainty on IBD selection efficiency also has a notable role.

## 1123 Sensitivities to NMO and oscillation parameters

1124 JUNO will start taking data in 2025. During the months and years to come, oscillation analyses will  
 1125 naturally be optimized regularly. What we described here represent the state of the art mid 2024,  
 1126 and was used for the sensitivity studies published in [61, 90] and are presented in table 2.7

	Central Value	PDG 2020	100 days	6 years	20 years
$\Delta m_{31}^2 (\times 10^{-3} \text{eV}^2)$	2.5283	$\pm 0.034$ (1.3%)	$\pm 0.021$ (0.8%)	$\pm 0.0047$ (0.2%)	$\pm 0.0029$ (0.1%)
$\Delta m_{21}^2 (\times 10^{-3} \text{eV}^2)$	7.53	$\pm 0.18$ (2.4%)	$\pm 0.074$ (1.0%)	$\pm 0.024$ (0.3%)	$\pm 0.017$ (0.2%)
$\sin^2 \theta_{12}$	0.307	$\pm 0.013$ (4.2%)	$\pm 0.0058$ (1.9%)	$\pm 0.0016$ (0.5%)	$\pm 0.0010$ (0.3%)
$\sin^2 \theta_{13}$	0.0218	$\pm 0.0007$ (3.2%)	$\pm 0.010$ (47.9%)	$\pm 0.0026$ (12.1%)	$\pm 0.0016$ (7.3%)

TABLE 2.7 – A summary of precision levels for the oscillation parameters. The reference value (PDG 2020 [92]) is compared with 100 days, 6 years and 20 years of JUNO data taking.

### 1127 Asimov studies

1128 To study the behavior and performance of fit procedures with enough realism, one should perform  
 1129 fits to a large number of toy spectra, generated with a number events equal to what one expects in  
 1130 real data, for the given exposure under consideration. This allows to study the impact of realistic  
 1131 statistical fluctuations. This is, however, time consuming, since thousands of spectra have to be  
 1132 generated and fitted.

1133 When subtle details are not crucial, another approach is possible to estimate sensitivities to the NMO  
 1134 and oscillation parameters, as well as (for instance) to verify the technical implementation of fitter  
 1135 (as we will do in Chapter 7 for the implementation of the joint fit). It consists on generating only 1  
 1136 pseudo-data sample, where the content of each bin  $D^i$  is set to the predicted value  $T^i$ , computed with  
 1137 a reasonable choice for the values of the model parameters (for instance, with the recent PDG values  
 1138 for the oscillation parameters). This is equivalent to a spectrum with fluctuations. It provides valid  
 1139 sensitivities if the expected statistics in the real data sample is high enough in each bin to assume a  
 1140 gaussian behavior.

### 1141 2.7.4 Versions of the fit used in this thesis

1142 In Chapter 7, we'll study the potential of a particular application of Dual Calorimetry, call "Dual  
 1143 Calorimetry with neutrino oscillation." This approach require to perform fits to the  $E^{vis}$  spectrum  
 1144 reconstructed with the LPMT system, with the SPMT system, and a joint fit to both spectra.

1145 In the two former cases, the PearsonV  $\chi^2$  introduced above will be used. In the latter case, it will  
 1146 be extended in the following way : The  $D$  data vector now possess 820 elements. Indeed, the fit is  
 1147 performed to a joint spectrum, where the LPMT spectrum is juxtaposed with the SPMT spectrum  
 1148 (see Figure 2.16).

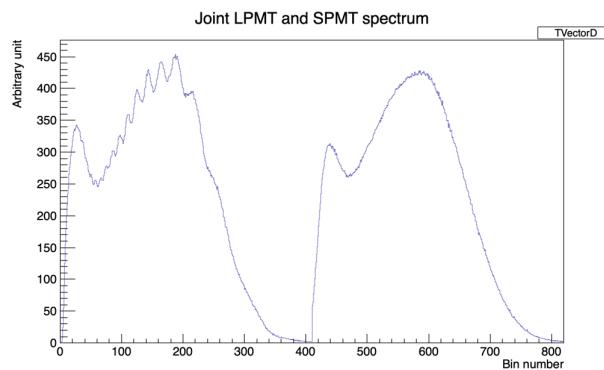


FIGURE 2.16 – Illustration of the spectrum considered when joint fitting

1149 The prediction vector  $T(\theta^d, \eta)$  is naturally extended in the same way. Its components 1 to 410  
 1150 predict the number of entries in the LPMT part of the LPMT+SPMT joint spectrum, while its com-  
 1151 ponents from 411 to 820 predict the contents of the SPMT part. Note that the list of oscillation  
 1152 parameters in  $T_{411}(\theta^d, \eta)$  to  $T_{820}(\theta^d, \eta)$  is the same as usual. However,  $T_1(\theta^d, \eta)$  to  $T_{410}(\theta^d, \eta)$  2  
 1153 additional parameters,  $\delta \sin^2(2\theta_{12})$  and  $\delta \Delta m_{21}^2$ , are added to the corresponding oscillation parameters  
 1154 to account for a potential unexpected problem in the LPMT reconstruction or calibration.

1155 In the case of this joint fit, the covariance matrix  $V$  is extended to a  $(820 \times 820)$  matrix. It is a central  
 1156 element of this study, as will be explained in Chapter 7, since the LPMT and SPMT data spectrum  
 1157 are correlated, even at the statistical level. The determination of this matrix will be an important and  
 1158 original point.

1159 Fits will be performed to an histogram spectrum defined over the 0.8-9 MeV range, with a flat binning  
 1160 (20 keV wide bins), often restricted to the 335 lowest  $E^{vis}$  bins.

1161 In this Section 2.7, we have provided a theoretical description of the fit procedures developed at  
 1162 JUNO. Software frameworks are necessary to use them in practice. The framework developed at  
 1163 Subatech will be described in Chapter 7.

## 1164 2.7.5 Physics results

1165 The oscillation parameters are directly extracted from the minimization procedure and the error can  
 1166 be estimated directly from the procedure. For the NMO, the data are fitted under the two assumption  
 1167 of NO and IO. The difference in  $\chi^2$  give us the preferred ordering and the significance of our test.  
 1168 Latest studies show that the precision on oscillation parameters after six year of data taking will be  
 1169 of 0.2%, 0.3%, 0.5% and 12.1% for  $\Delta m_{31}^2$ ,  $\Delta m_{21}^2$ ,  $\sin^2 \theta_{12}$  and  $\sin^2 \theta_{13}$  respectively [61]. The expected  
 1170 sensitivity to mass ordering is  $3\sigma$  after 6.5 years [53].

## 1171 2.8 Summary

1172 JUNO is one the biggest new generation neutrino experiment. Its goal, the measurements of oscil-  
 1173 lation parameters with unprecedented precision and an NMO preference at the 3 sigma confidence  
 1174 level, needs an in depth knowledge and understanding of the detector and the physics at hand. The  
 1175 characterisation and calibration of the detector are of the utmost importance and the understanding  
 1176 of the detector response in its resolution and bias is capital to be able to correctly carry the high  
 1177 precision physics analysis of the neutrino oscillation.

1178 In this thesis, I explore the usage of data-driven reconstruction methods to validate and optimize the  
 1179 reconstruction of IBD events in JUNO in the chapters 4, 5 and 6 and the usage of the dual calorimetry  
 1180 in the detection of possible mis-modelisation in the theoretical spectrum 7.

<sup>1181</sup> **Chapter 3**

<sup>1182</sup> **Introduction to the reconstruction  
methods and algorithms used in this  
thesis**

<sup>1183</sup>

<sup>1184</sup>

<sup>1185</sup> “I have the shape of a human being and organs equivalent to those of a human being. My organs, in fact, are identical to some of those in a prosthetized human being. I have contributed artistically, literally, and scientifically to human culture as much as any human being now alive. What more can one ask?”

*Isaac Asimov, The Complete Robot*

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<sup>1204</sup> <sup>1205</sup> Machine Learning (ML) and more specifically Neural Network (NN) are families of data-driven algorithms. They are used in a wide variety of domains including natural language processing, computer vision, speech recognition and, the subject of this thesis, scientific studies.

<sup>1206</sup> Machine learning models aim to learn underlying patterns from finite datasets in order to make <sup>1207</sup> general predictions or classifications. For example, in our case, it could be an algorithm that would <sup>1208</sup> differentiate the nature of a particle interacting in the liquid scintillator, between a positron and an <sup>1209</sup> electron, based on the readout charge and time ( $Q, t$ ) of the 17612 LPMT of the JUNO experiment. <sup>1210</sup> During a first training phase, it would learn the discriminative features between the two in the 35224-dimensional <sup>1211</sup> charge and time distribution, built from samples of  $e^+$  and  $e^-$  events.

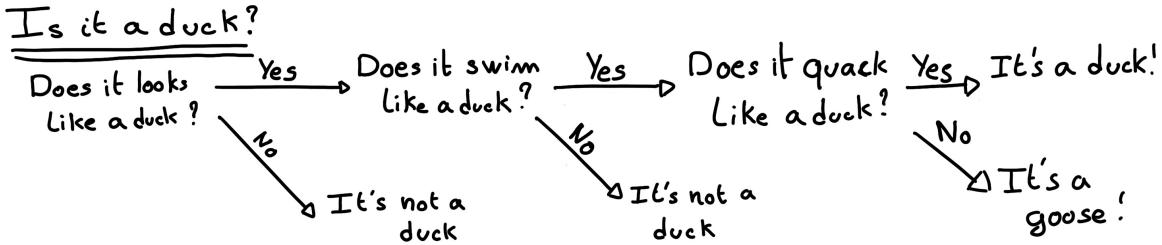


FIGURE 3.1 – Example of a BDT that determine if the given object is a duck

1215 It extracts essential features from a highly complex and multi-dimensional dataset that describe the  
 1216 physical interactions: a three body energy deposition (the positron and two annihilation gammas)  
 1217 and the single deposit from an electron.

1218 Ideally, the algorithm would learn to recognize those informations on its own, regardless of the input  
 1219 size and complexity. In practice, however, these algorithms are guided by human design through  
 1220 their architectures and training conditions. We can still hope that they can use more thoroughly the  
 1221 detector informations while traditional methods are often subject to assumptions or simplifications  
 1222 to make the task easier (see for instance the algorithm in Section 3.3).

1223 The role of machine learning algorithms has expanded rapidly in the past decade, either as the  
 1224 main or secondary algorithm for a wide variety of tasks: event reconstruction, event classification,  
 1225 waveform reconstruction and so on. In particular in domains where the underlying physic and  
 1226 detector processes are complex and highly dimensional, and when large amount of data must be  
 1227 processed quickly.

1228 This chapter present an overview of the different kind of machine learning methods and neural  
 1229 networks that will be discussed in this thesis, and the state of the art of the reconstructions methods  
 1230 in JUNO our ML algorithms will be compared to.

### 1231 3.1 Core concepts in machine learning and neural networks

1232 In this section, we discuss the core concepts in machine learning that will be used thorough this  
 1233 thesis. We place particular emphasis on Neural Networks, as it's the family of the algorithms  
 1234 described in chapters 4, 5 and 6.

#### 1235 3.1.1 Boosted Decision Tree (BDT)

1236 One of the most classic machine learning algorithm used in particle physics is Boosted Decision Tree  
 1237 (BDT) [93] (or more recently Gradient Boosting Machine [94]).

1238 BDTs operate by making a series of decisions based on a set of input features, with each decision  
 1239 represented as a node in the tree. Each decision point, or node, takes its decision based on a set of  
 1240 trainable parameters leading to a subtree of decisions. The process is repeated until it reach the final  
 1241 node, yielding the prediction. A simplistic example is given in Figure 3.1.

1242 The training procedure follows a reward-based approach where the algorithm predictions are com-  
 1243 pared to the true outcomes. During the training phase the prediction of the BDT is compared to a  
 1244 known truth about the data. The score is then used to backpropagate corrections to the parameters  
 1245 of the tree. Modern BDT use gradient boosting where the gradient of the loss is calculated for each  
 1246 of the BDT parameters. Following the gradient descent, we can reach the, hopefully, global minima  
 1247 of the loss for our set of parameters.

### 3.1.2 Artificial Neural Network (NN)

One of the modern ML family is the Neural Network, historical name as their design was inspired by the behaviour of biological neurons in the brain. As schematized in Figure 3.2, the input, output

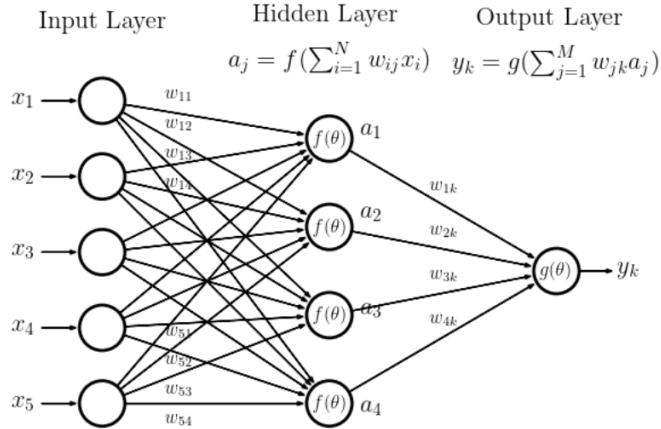


FIGURE 3.2 – Schema of a simple neural network

and steps inside the NN is described as neuron *layers*. The neurons of the layers take as input a set of values from the preceding layer, here the  $a_i$  takes every informations of the  $x_i$  input layer, and aggregate those values following learnable *parameters*  $w_{ij}$ . In the example in Figure 3.2, fully connected layers are used, meaning that each neuron in one layer is connected to every neuron in the previous layer.

The aggregation procedure is core of defining the architecture of the NN. The different architectures used in this thesis will be discussed in Section 3.2. The process is repeated until reaching the output layer.

For example, let's take the network in Figure 3.2 and say that  $a_1$ ,  $a_2$  and  $a_3$  are the neurons of the output layer. We try to produce a vertex reconstruction algorithm that will approach the charge barycentre. Let's limit the input  $x_i$  to the charge of the  $i$ th PMT, one of the solution is to aggregate on  $a_1$  the  $x$  coordinate of the barycenter. The network would thus adapt the  $w_{i1}$  parameters so they correspond to the  $x$  coordinates of the  $i$ th PMT. Same for the  $y$  and  $z$  coordinate on  $a_2$  and  $a_3$  respectively.

The layers used in the example above are designated as *Fully connected* layers, where every neurons of the layer is connected to the every neurons of the preceding layer. The layer can be expressed using the Einstein summation and in bold the learnable parameters

$$O_j = I_i + \mathbf{W}_j^i \quad (3.1)$$

where  $O_j$  is the output neurons vector (the  $a_i$ ),  $I_i$  is the preceding layer neurons vector (the  $x_i$ ) and  $\mathbf{W}$  is the parameters, or weights, matrix (composed of the  $w_{ij}$ ). In practice, this fully connected layer is often adjoined a bias  $B$  and an *activation function*  $F$ .

$$I_j = F(I_i \mathbf{W}_j^i + \mathbf{B}_j) \quad (3.2)$$

This is the fundamental component of the Fully Connected Deep NN (FCDNN) family presented in Section 3.2.1.

This description of neural networks as layers introduce the principles of *depth* and *width*, the number

1274 of layers in the NN and the number of neurons in each layer respectively. Those quantities that not  
 1275 directly used for the computation of the results but describes the NN or its training are designated  
 1276 as *hyperparameters*.

1277 Now we just need to adapt the parameters so that this network learn that  $w_{ij}$  are the PMT coordinate.  
 1278 We describe the space produced by the parameters of the network as the *parameter phase space* or *latent*  
 1279 *space*. The optimization of the network and exploration of this phase space is done through training  
 1280 over a *training dataset* as described in next section.

### 1281 3.1.3 Training procedure

1282 To adapt the parameters we need an object that describe how well the network perform. This is  
 1283 the *loss* of our neural networks  $\mathcal{L}$ . In our barycenter example, it could be the distance between the  
 1284 reconstructed and real barycenter. Using this metric we can adjust the parameters of our network.

1285 Depending if we try to minimize or maximize it, it need to posses a minima or a maxima. For example  
 1286 when doing *regression*, i.e. produce a scalar result like the coordinates of a barycenter, a common loss  
 1287 is the Mean Square Error (MSE). Let  $i$  be our dataset, the  $N$  events considered for training,  $y_i$  be the  
 1288 target scalar, the barycenter positions of each events,  $x_i$  the input data, the charge vector, and  $f(x_i, \theta)$   
 1289 the result of the network. The network here is modelled by  $f$ , and its parameter  $\theta$

$$\mathcal{L} \equiv MSE = \frac{1}{N} \sum_i^N (y_i - f(x_i, \theta))^2 \quad (3.3)$$

1290 Another common loss function is the Mean Absolute Error (MAE)

$$\mathcal{L} \equiv MAE = \frac{1}{N} \sum_i^N |y_i - f(x_i, \theta)| \quad (3.4)$$

1291 We see that those loss function possess a minima when  $f(x_i, \theta) = y_i$ .

1292 Modern neural networks typically use gradient descent to optimize their parameters by minimizing  
 1293 the loss function. The gradient of the parameter  $w$ , designated in literature as  $\theta$ , with respect of the  
 1294 loss function  $\mathcal{L}$  is subtracted each optimisation step  $t$

$$\theta_{t+1} = \theta_t - \frac{\partial \mathcal{L}}{\partial \theta} \quad (3.5)$$

1295 This induce  $\mathcal{L}$  needs to be differentiable with respect to  $\theta$ , thus the layers and their activation func-  
 1296 tions also need to be differentiable. This simple gradient descent, designated as Stochastic Gradient  
 1297 Descent (SGD), can be extended with first and second order momentums like in the Adam optimizer  
 1298 [95]. More details about the optimizers can be found in Section 3.1.3.

### 1299 Training lifecycle

1300 The training process of neural networks can vary depending on the application and dataset, but in  
 1301 this thesis, we follow a standard approach. As shown in Fig. 3.3, training is organized into *epochs*,  
 1302 each of which consists of several *steps*. During each step, the neural network optimizes its parameters  
 1303 using a *batch*, a subset of the entire training dataset.

1304 The ideal batch size, meaning the number of events in each batch, would encompass the entire  
 1305 dataset to avoid bias introduced by sub-sample specificity. However, in large-scale experiments  
 1306 like JUNO, the batch size is often constrained by memory limitations due to the massive volume of

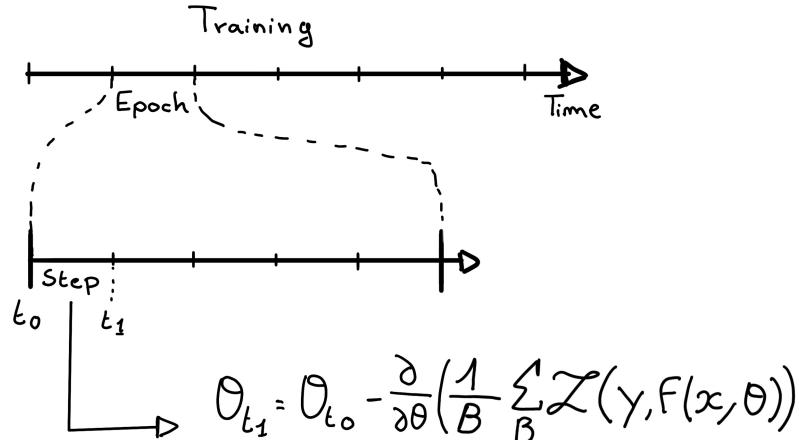


FIGURE 3.3 – Illustration of the training lifecycle

1307 data generated by the photomultiplier tubes (PMTs). Balancing batch size with memory capacity is  
 1308 crucial to ensure efficient and accurate training.

1309 At the end of each epoch, the neural network is evaluated on a validation dataset, which is not  
 1310 used during training. This dataset serves as a reference to assess the network's performance and to  
 1311 monitor for signs of overfitting. In JUNO, this is critical because the model needs to generalize well  
 1312 to unseen experimental data and avoid overfitting to noise in the training set (see Section 3.1.4).

1313 Hyperparameters that can be optimized during the training can be optimized at each epoch, for  
 1314 example the learning rate, or each step, the optimizer momentum for example.

1315 There is not really a typical number of epochs or steps for the training. The number steps can be  
 1316 defined such as in one epoch, the NN see the entirety of the dataset but the number of steps and  
 1317 epochs are hyperparameters that are optimized over the each subsequent training. We adjust them  
 1318 by looking at the loss evolution profile over time.

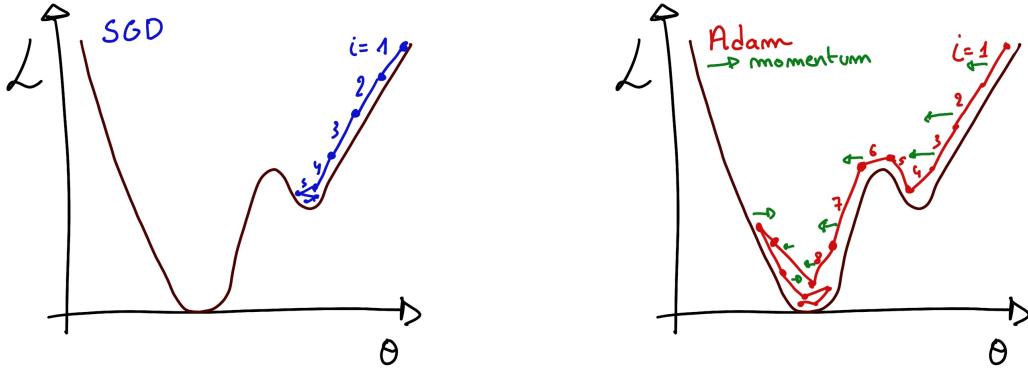
1319 Most training are started with a fixed number of epochs, i.e. from what we've seen from precedent  
 1320 training, the network stop learning, the loss is constant, after  $N$  epoch so we run the training for  
 1321  $N + \delta$  epochs to see if the modification brings improvements to the loss profile. We can implement  
 1322 *early stopping policies* to halt training if certain conditions are met, such as a sudden increase in loss  
 1323 or when the loss plateaus. However, for the JUNO experiment, where training time is not a strict  
 1324 limitation, early stopping is less critical, though it may still be useful to prevent overfitting in some  
 1325 cases

### 1326 The optimizer

1327 As briefly introduced at the beginning of this section, the parameters of the neural network are  
 1328 optimized using the gradient descent method. We compute the gradient of the mean loss over the  
 1329 batch with respect of each parameters and we update the parameters in accord to minimize the loss.  
 1330 The gradient is computed backward from the loss up to the first layer parameters using the chain  
 1331 rule, in this case with only one parameter at each step for simplicity:

$$\frac{\partial \mathcal{L}}{\partial \theta_1} = \frac{\partial \theta_2}{\partial \theta_1} \frac{\partial \mathcal{L}}{\partial \theta_2} = \frac{\partial \theta_2}{\partial \theta_1} \frac{\partial \theta_3}{\partial \theta_2} \frac{\partial \mathcal{L}}{\partial \theta_3} = \frac{\partial \theta_2}{\partial \theta_1} \prod_{i=2}^{N-1} \frac{\partial \theta_{i+1}}{\partial \theta_i} \frac{\partial \mathcal{L}}{\partial \theta_N} \quad (3.6)$$

1332 where  $\theta$  is a parameter,  $i$  is the layer index. We see here that the gradient of the first layer is  
 1333 dependent of the gradient of all the following layers. Because the only value known at the start



(A) Illustration of SGD falling into a local minima

(B) Illustration of the Adam momentum allowing it to overcome local minima

FIGURE 3.4

of the optimization procedure is  $\mathcal{L}$  we compute  $\frac{\partial \mathcal{L}}{\partial \theta_N}$  then,  $\frac{\partial \theta_N}{\partial \theta_{N-1}}$ , etc... This is called the *backward propagation*.

This update of the parameters is done following an optimizer policy. Those optimizers depends on hyperparameters. The ones used in this thesis are:

1. Stochastic Gradient Descent (SGD). A simple but widely used optimizer that relies on one key hyperparameter, the learning rate (LR) /  $\lambda$ . It update each step the parameters  $\theta$  following

$$\theta_{t+1} = \theta_t - \lambda \left. \frac{\partial \mathcal{L}}{\partial \theta} \right|_{\theta_t} \quad (3.7)$$

where  $t$  is the step index. It is a powerful optimizer but is very sensible to local minima of the loss in the parameters phase space as illustrated in Figure 3.4a.

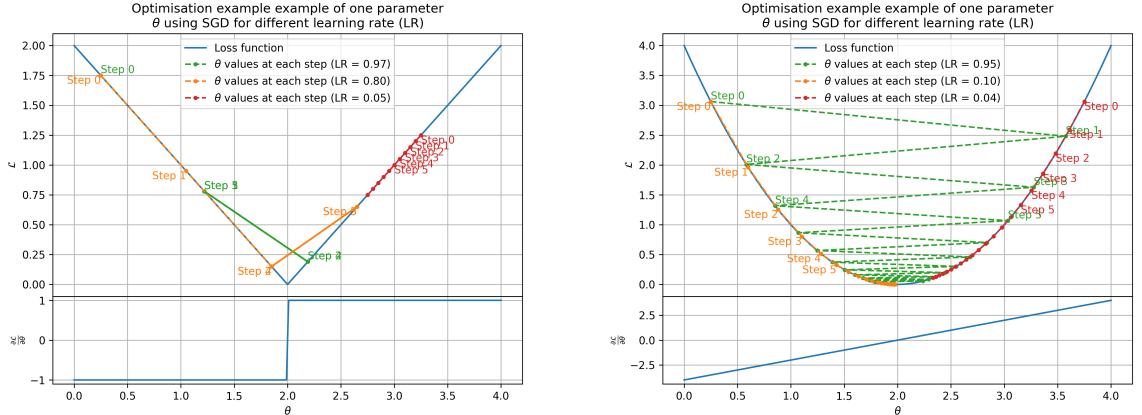
2. Adam Optimizer [95]. The concept is, in short, to have and SGD but with momentum. Adam possess two momentum  $m(\beta_1)$  and  $v(\beta_2)$  which are respectively proportional to  $\frac{\partial \mathcal{L}}{\partial \theta}$  and  $(\frac{\partial \mathcal{L}}{\partial \theta})^2$ .  $\beta_1$  and  $\beta_2$  are hyperparameters that dictate the moment update at each optimization step. The parameters are then upgraded following

$$m_{t+1} = \beta_1 m_t + (1 - \beta_1) \frac{\partial \mathcal{L}}{\partial \theta} \quad (3.8)$$

$$v_{t+1} = \beta_2 v_t + (1 - \beta_2) \left( \frac{\partial \mathcal{L}}{\partial \theta} \right)^2 \quad (3.9)$$

$$\theta_{t+1} = \theta_t - \lambda \frac{m_{t+1}}{\sqrt{v_{t+1}} + \epsilon} \quad (3.10)$$

where  $\epsilon$  is a small number to prevent divergence when  $v$  is close to 0. These momentums allow to overcome small local minima in the parameters phase. Imagine ball going down a slope as illustrated in 3.4a, if you ignore the stored momentum you get SGD and get stuck as on the left plot. Now if you consider the momentum you get over the hill and end up in the global minima.



(A) Illustration of the SGD optimizer on one parameter  $\theta$  on the MAE Loss. We see here that it has trouble reaching the minima due to the gradient being constant.

(B) Illustration of the SGD optimizer on one parameter  $\theta$  on the MSE Loss. We see two different behavior: A smooth one (orange and red) when the LR is small enough and a more chaotic one when the LR is too high.

FIGURE 3.5 – Illustration of the SGD optimizer. In blue is the value of the loss function, orange, green and red are the path taken by the optimized parameter during the training for different LR.

### 1347 Learning Rate (LR) Schedules

1348 The learning rate plays a crucial role in determining how fast or slow the model converges. If the  
 1349 learning rate is too high (Fig. 3.5a), the model may skip over the optimal solution, whereas a low  
 1350 learning rate (Fig. 3.5b) can slow down the convergence process, leading to inefficient training. To  
 1351 address this, learning rate schedulers are employed.

1352 Using a learning rate scheduler allows the optimizer to take larger steps in the early stages of training,  
 1353 where rapid learning is beneficial, and progressively smaller steps as the model approaches convergence.  
 1354 This strategy is especially useful in JUNO, where early learning from noisy data may require  
 1355 coarse adjustments, but fine-tuning is needed later to accurately capture subtle event characteristics.

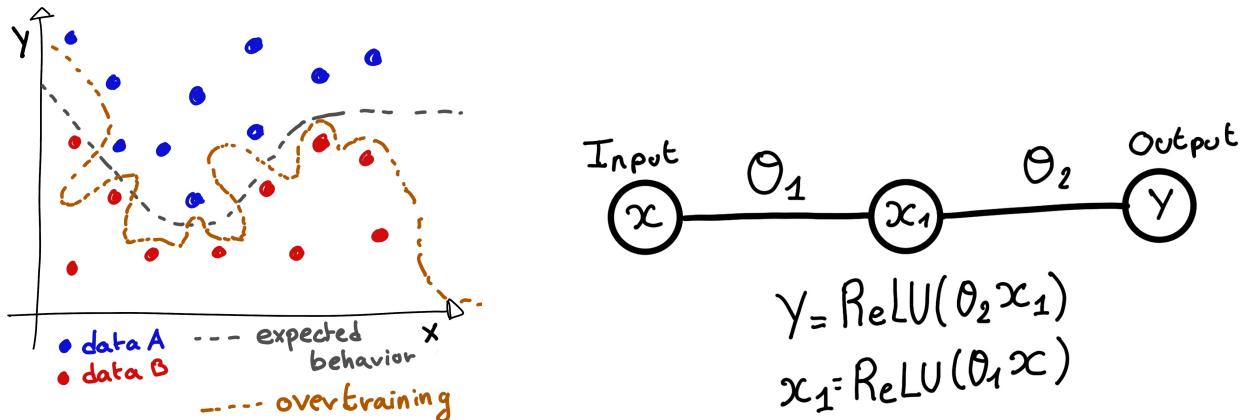
1356 Another policy that is often used is the save of the best model. In some situations, the loss value after  
 1357 each epoch will strongly oscillate or even worsen. This policy allows us to keep the best version  
 1358 of the model attained during the training phase.

### 1359 3.1.4 Potential pitfalls

1360 Apart from being stuck in local minima, there are also other behaviors and effects we want to prevent  
 1361 during training.

#### 1362 Overtraining

1363 Overfitting occurs when a neural network memorizes specific details or noise from the training  
 1364 dataset rather than learning a general representation of the underlying data. This is common when  
 1365 the training dataset is small relative to the number of parameters in the network or when the dataset  
 1366 contains specific features that do not generalize well to unseen data. Additionally, training the



(A) Illustration of overtraining. The task at hand is to determine depending on two input variable  $x$  and  $y$  if the data belong to the dataset  $A$  or the dataset  $B$ . The expected boundary between the two dataset is represented in grey. A possible boundary learnt by overtraining is represented in brown.

(B) Illustration of a very simple NN

FIGURE 3.6

network for too many epochs can exacerbate this issue. Figure 3.6a illustrates the impact of overfitting, where the model fits the training data too closely, compromising its ability to generalize. To detect overfitting, techniques like monitoring the validation loss, early stopping, or employing cross-validation can be employed. In JUNO's context, managing overfitting is critical due to the large volume of data generated by the photomultiplier tubes (PMTs), which may include noise or other artifacts.

Overtraining can be fought in multiple ways, for example:

- **More data.** By having more data in the training dataset, the network will not be able to learn the specificities of every data.
- **Less parameters.** By reducing the number of parameters, we reduce the computing and learning capacities of the network. This will force it to fallback to generalist behaviours.
- **Dropout.** This technique implies to randomly set some neurons to 0, i.e. cutting the relation between two neurons in a layer. By doing this, we force the network to allocate more of its parameter to the features learning, preventing those parameters to be used for overtraining.
- **Early stopping.** During the training we monitor the network performance over a validation dataset. The network does not train on this dataset and thus cannot learn its specificities. If the loss on the training dataset diverge too much from the loss on the validation dataset, we can stop the training earlier to prevent it from overtraining.

### 1385 Gradient vanishing

1386 Gradient vanishing is the effect of the gradient being so small for the early layers that the parameters  
 1387 are barely updated after each step. This cause the network to be unable to converge to the minima.

1388 This comes from the way the gradient descent is calculated. Imagine a simple network composed of  
 1389 three fully connected layers: the input layer, a intermediate layer and the output layer. Let  $L$  be the  
 1390 loss,  $\theta_1$  the parameter between the input and the intermediate layer and  $\theta_2$  the parameter between  
 1391 the intermediate and output layer. This network is schematized in Figure 3.6b.

1392 The gradient for  $\theta_1$  will be computed using the chain rule presented in equation 3.6. Because  $\theta_1$   
 1393 depends on  $\theta_2$ , if the gradient of  $\theta_2$  is small, so will be the gradient of  $\theta_1$ . Now if we would have  
 1394 much more layer, we can see how the subsequent multiplication of small gradients would lead to  
 1395 very small update of the parameters thus “*vanishing gradient*”.

1396 Multiple actions can be taken to prevent this effect such as:

- 1397 — **Batch normalization:** In this case we apply a normalization layer that will normalize the data.  
 1398 It means that we transform the input variable  $X$  into a variable  $D$  which distribution follow  
 1399  $\langle D \rangle = 0$  and  $\sigma_D = 1$ . This helps the parameters of the network to maintain an appropriate  
 1400 scale.
- 1401 — **Residual Network (ResNet) [96]:** Residual network is a technique for neural network in which,  
 1402 instead of just sequentially feeding the results of each layer to the next one, you compute a  
 1403 residual over the input data. This technique is illustrated in Figure 3.7. The reference [96] show  
 1404 empirical evidence of its relevance.

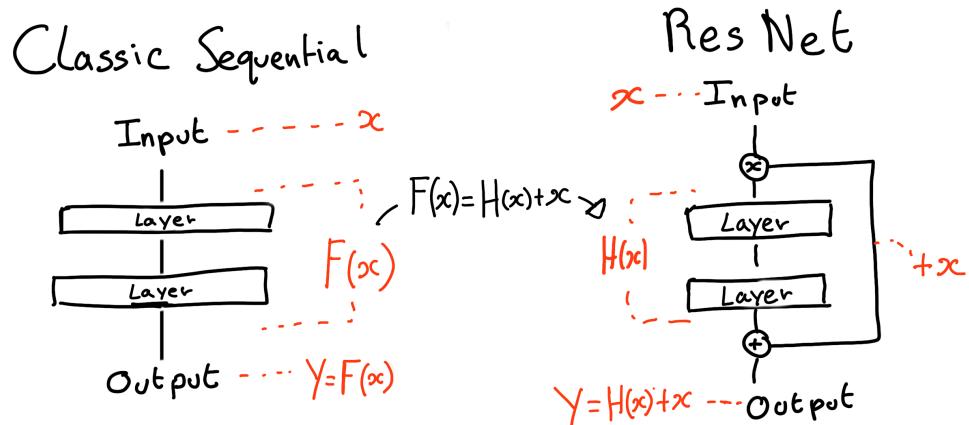


FIGURE 3.7 – Illustration of the ResNet framework

### 1405 Gradient explosion

Gradient explosion occurs when gradients grow exponentially during backpropagation, causing parameter values to increase dramatically. This is particularly problematic in deep networks where the product of large gradients across layers can lead to unstable updates. In practice, gradient explosion is often caused by large learning rates, poor weight initialization, or nonlinearities in the network. For illustration, consider that the loss dependency in  $\theta$  follow

$$\begin{aligned}\mathcal{L}(\theta) &= \frac{\theta^2}{2} + e^{4\theta} \\ \frac{\partial \mathcal{L}}{\partial \theta} &= \theta + 4e^{4\theta}\end{aligned}$$

1406 The explosion is illustrated in Figure 3.8 where we can see that the loss degrades with each step of  
 1407 optimization. In this illustration it is clear that reducing the learning rate suffice but this behaviour  
 1408 can happens in the middle of the training where the learning rate schedule does not permit reactivity.

1409 There exist solutions to prevent this explosions:

- 1410 — **Gradient clipping:** In this case we work on the gradient so that the norm of gradient vector  
 1411 does not exceed a certain threshold. In our illustration in Figure 3.8 the gradient for  $\theta > 0$   
 1412 could be clipped at 3 for example.

- 1413 — **Batch normalization:** For the same reasons as for gradient vanishing, normalizing the input  
 1414 data help reduce erratic behaviour.

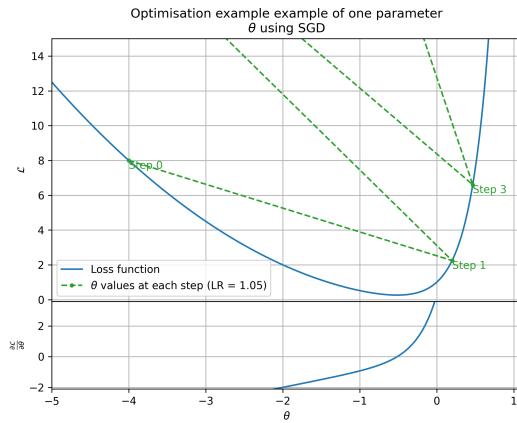


FIGURE 3.8 – Illustration of the gradient explosion. Here it can be solved with a lower learning rate but its not always the case.

## 1415 3.2 Neural networks architectures

### 1416 3.2.1 Fully Connected Deep Neural Network (FCDNN)

1417 In this thesis, FCDNN serves as a baseline architecture for comparison with more specialized models  
 1418 like CNNs (see Section 3.2.2) and GNNs (see section 3.2.3), which are better suited to structured or  
 1419 graph-based data. However, FCDNNs are still useful when modeling highly abstract relationships,  
 1420 such as aggregating features from the JUNO PMTs. While they are powerful, their main drawback  
 1421 lies in their inefficiency when dealing with high-dimensional or spatially structured data, which  
 1422 will be addressed with convolutional architectures. This architecture is the stack of multiple fully  
 1423 connected layers as presented in the Figure 3.9a. Most of the time, the classic ReLU function

$$\text{ReLU}(x) = \begin{cases} x & \text{if } x \geq 0 \\ 0 & \text{otherwise} \end{cases} \quad (3.11)$$

1424 is used as activation function. Prelu and Sigmoid are also popular choices:

$$\text{Sigmoid}(x) = \frac{1}{1 + e^{-x}} \quad (3.12) \quad \text{PReLU}(x) = \begin{cases} x & \text{if } x \geq 0 \\ \alpha x & \text{otherwise} \end{cases} \quad (3.13)$$

1426 The reasoning behind ReLU and PReLU is that with enough of them, you can mimic any continuous  
 1427 function as illustrated in Figure 3.9b. Sigmoid is more used in case of classification, its behavior  
 1428 going hand in hand with the Cross Entropy loss function used in classification problems.

1429 Due to its simplicity, FCDNN are also used as basic pieces for more complex architectures such as  
 1430 the CNN and GNN that will be presented in the next sections.

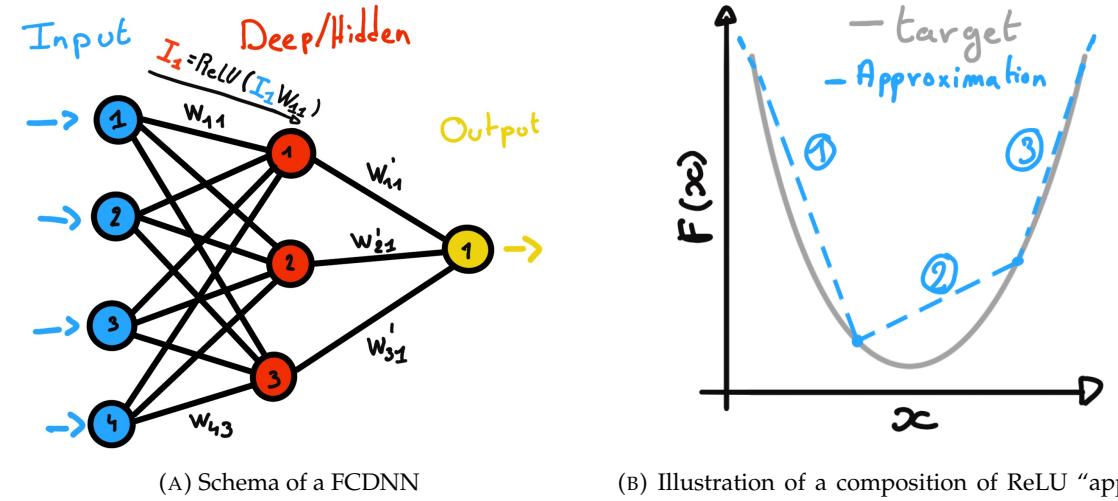


FIGURE 3.9

### 3.2.2 Convolutional Neural Network (CNN)

It's not trivial to describe in text the principles of Convolutional Neural Network (CNN) and how they works. We try a general description below followed by a step by step description of a concrete example.

Convolutional Neural Networks are a family of neural networks that use discrete convolution filters, as illustrated in an example in Figure 3.10, to process the input data, often images. They are commonly used in image recognition [97] for classification or regression problematics. Concretely, you multiply element-wise a portion of the input data, in the case of an image, a small part of the image, with a kernel of same dimension. In Figure 3.10, we multiply the  $3 \times 3$  pixels sub-image with the  $3 \times 3$  kernel.

Their filters scan the input data, highlighting patterns of interest, this scanning procedure making them translation-invariant. In the concrete case of Figure 3.10, for each pixel of the input image, we group it with the 8 neighbours pixel and produce a new pixel that correspond to the output image. For the pixel on the edges that do not have neighbours, we either create "imaginary" pixel with the value 0 or we just ignore them. If we ignore them, the output image will posses fewer pixels than the input image. We see that the operation do not care where is the pattern of interest in the images, the filter output will be *invariant* whatever *translation* is applied to the image.

This invariance mean that they are capable of detecting oriented features independently of their location on the image. These filters scan the input, highlighting important features like edges or textures, which in JUNO's case could represent spatial correlations in the timing and charge data across the detector. As the network goes deeper, it can capture more complex and abstract features, making it ideal for detecting nuanced particle interactions. Again taking 3.10 as an example, with only the 9 parameters composing the kernel, we can highlight the contour of the duck by looking at the "yellowness" of the pixels.

The learning parameters of CNNs are the kernels components, the network thus learn the optimal filters to extract the desired features.

The convolution layers are commonly chained [98], reducing the input dimension while increasing the number of filters. The idea behind is that the first layers will process local informations and the latest layers will process more global informations, as the latest convolution filters will process

the results of the preceding that themself have processed local information. To try to preserve the amount of information, we tend to grow the numbers of filters for each division of the input data. The results of the convolution filters is commonly then flattened and feed to a smaller FCDNN which will process the filters results to yield the desired output.

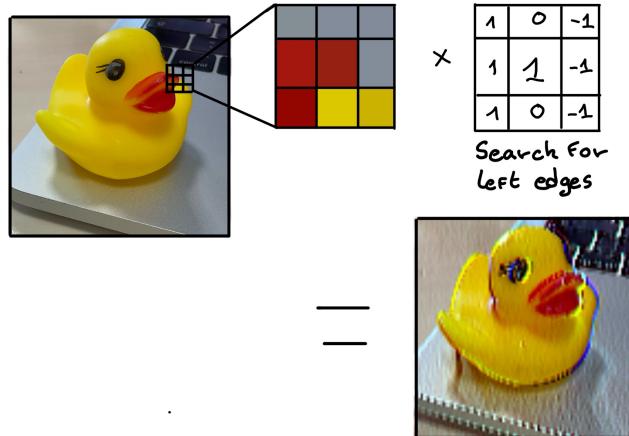


FIGURE 3.10 – Illustration of the effect of a convolution filter. Here we apply a filter with the aim do detect left edges. We see in the resulting image that the left edges of the duck are bright yellow where the right edges are dark blue indicating the contour of the object. The convolution was calculated using [99].

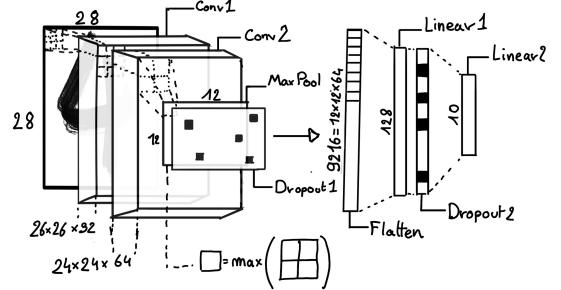
As an example, let's take the Pytorch [100] example for the MNIST [101], a dataset of black and white images of handwritten digits. Those images are  $28 \times 28$  pixels with only one channel corresponding to the grey level of the pixel. Example of images from this dataset are presented in Figure 3.11a

A schema of the CNN used in the Pytorch example is presented in Figure 3.11b. Using this schema as a reference, the trained network is made of:

1. A convolutional layer of  $(3 \times 3)$  filters yielding 32 channels. A bias parameter is applied to each channel for a total of  $(32 \cdot (3 \times 3) + 32) = 320$  parameters. The resulting image is  $(26 \times 26 \times 32)$  (26 per 26 pixels with 32 channels). The ReLU activation function is applied to each pixel.
2. A second convolutional layer of  $(3 \times 3)$  filters yielding 64 channels. This channel also posses a bias parameter for a total of  $(64 \cdot (3 \times 3) + 64) = 640$  parameters. Resulting image is  $(24 \times 24 \times 64)$ . This channel also apply a ReLU activation function.
3. Then comes a  $(2 \times 2)$  max pool layer with a stride of 1 meaning that for each channel the max value of pixels in a  $(2 \times 2)$  block is condensed in a single resulting pixel. The resulting image is  $(12 \times 12 \times 64)$ .
4. This image goes through a dropout layer which will set the pixel to 0 with a probability of 0.25. This help prevent overtraining the neural network (see Section 3.1.4 for more details).
5. The data is the flattened i.e. condensed into a vector of  $(12 \times 12 \times 64) = 9216$  values.
6. Then comes a fully connected linear layer (Eq. 3.2) with a ReLU activation that output 128 feature. It needs  $(9216 \cdot 128) + 128 = 1'179'776$  parameters.
7. This 128 item vector goes through another dropout layer with a probability of 0.5
8. The vector is then transformed through a linear layer with ReLU activation. It output 10 values, one for each digit class  $(0, 1, 2, \dots, 9)$ . It need  $(128 \cdot 10) + 128 = 1408$  parameters.
9. Finally the 10 values are normalized using a log softmax function  $\text{LogSoftmax}(x_i) = \log \left( \frac{\exp(x_i)}{\sum_j \exp(x_j)} \right)$ . Each of those values are the probability of the input image to be a certain digit.



(A) Example of images in the MNIST dataset



(B) Schema of the CNN used in Pytorch example to process the MNIST dataset

FIGURE 3.11

1488 The final network needs 1'182'144 parameters or, if we consider each parameters to be a double  
 1489 precision floating point, 9.45 MB of data. To gives a order of magnitude, such neural network is  
 1490 considered "simple", train in a matter of minutes on T4 GPU [102] (14 epochs) and reach an accuracy  
 1491 in its prediction of 99%.

### 1492 3.2.3 Graph Neural Network (GNN)

1493 In GNNs, data is represented as nodes and edges in a graph, which allows us to model the JUNO  
 1494 detector as a network of PMTs, where each PMT is a node and the edges represent relationships  
 1495 such as spatial distance or timing correlations between PMTs. This flexibility enables GNNs to  
 1496 capture complex interactions across the detector geometry that would be difficult to represent with a  
 1497 CNN. Furthermore, GNNs excel at processing non-Euclidean data, making them a natural fit for the  
 1498 irregular layout of the PMTs in JUNO. In this thesis, GNNs are applied to model the spatial and tem-  
 1499 poral relationships between PMTs, enabling more precise event classification and reconstruction. By  
 1500 leveraging the message-passing framework, the GNN can aggregate information from neighboring  
 1501 PMTs, allowing it to detect subtle patterns in the detector's data.

1502 To get deeper in details, we have seen in the previous section, the CNNs are powerful for image  
 1503 processing, and more generally any data that can be expressed as a regular, discrete space and from  
 1504 which the information reside in the dispersion in this space. For an image, the edges of an object  
 1505 and how they assemble. A red square, straight edges with a sharp angle between them, is much less  
 1506 representative of a duck than an yellow sphere, round edges without sharp angles.

1507 This "image" projection is not fitted for every problematics. The signals produced by a detector does  
 1508 not always have the properties of images. In the case of JUNO for example, we can create an image  
 1509 of two channels, one for the charge  $Q$  and one for the timing  $t$  but this image should be spheric.  
 1510 Furthermore JUNO is by nature inhomogeneous, using two different systems : The LPMT and the  
 1511 SPMT. Those two systems have different regime, and thus should be processed differently. We could  
 1512 imagine images with four channels, two for the LPMT and two for the SPMT, or even a branched  
 1513 CNN with one convolution branch for the LPMT and another one for the SPMT. Anyway, the CNN  
 1514 will need to combine the two systems.

1515 To get around the restrictions of data representation imposed by CNNs, we can use the more flexible  
 1516 *graph* representation. A graph  $G(\mathcal{N}, \mathcal{E})$  is composed of vertex or node  $n \in \mathcal{N}$  and edges  $e \in \mathcal{E}$ . The  
 1517 edges are associated to two nodes  $(u, v) \in \mathcal{N}^2$ , "connecting" them. The node and the edges can hold  
 1518 features, commonly represented as vector  $n \in \mathbb{R}^{k_n}$ ,  $e \in \mathbb{R}^{k_e}$  with  $k_n$  and  $k_e$  the number of features on  
 1519 the nodes and edges respectively. We can thus define a graph using two tensors  $A_e^{ij}$  the adjacency

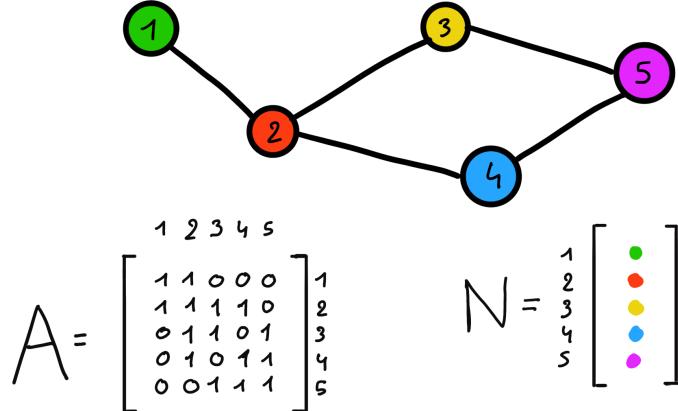


FIGURE 3.12 – Illustration of a graph and its tensor representation.

1520 tensor that hold the features  $\epsilon \in [0, k_e]$  of the edge connecting the node  $i$  and  $j$  and the tensor  $N_v^i$  that  
1521 hold the features  $v \in [0, k_n]$  of a node  $i$ .

1522 More figuratively, using the example in Figure 3.12, we have a graph of 5 nodes with a color as  
1523 feature. The edges have no features, we thus encode their existences as 0 or 1. In a realistic examples  
1524 as JUNO we could represent each PMTs as nodes and the edges between them as their relation such  
1525 as distance, timing difference, etc... There no strict rules about what is a node or how they should be  
1526 linked together. This abstraction allow us to represent virtually any type of detector of any geometry.

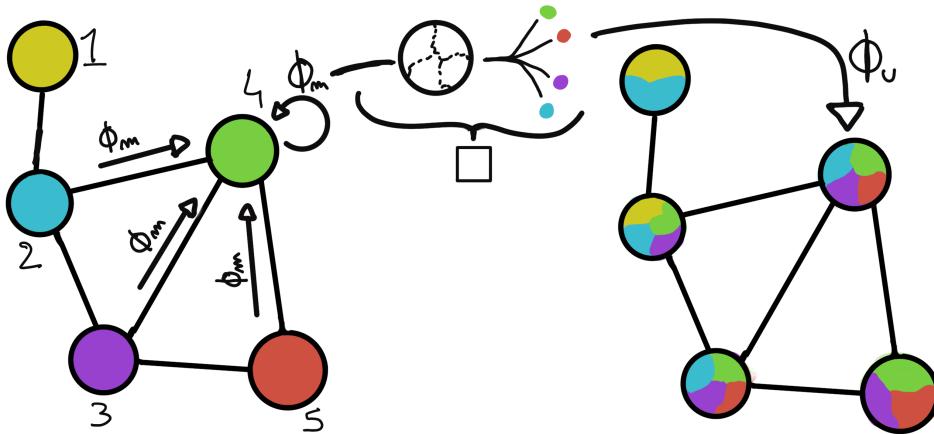


FIGURE 3.13 – Illustration of the message passing algorithm. The detailed explanation can be found in Section 3.2.3

1527 To process such object we need specific machine learning algorithms we call Graph neural network.  
1528 To efficiently manipulate graph we need to structurally encode their property in the neural network  
1529 computing architecture: each node is equivalent (as opposite to ordered data in a vector), each node  
1530 has a set of neighbours, ... One of this method is the message passing algorithm presented historically  
1531 in "Neural Message Passing for Quantum Chemistry" [103]. In this algorithm, with each layer of  
1532 message passing a new set of features is computed for each node following

$$n_i^{k+1} = \phi_u(n_i^k, \square_j \phi_m(n_i^k, n_j^k, e_{ij}^k)); n_j \in \mathcal{N}_i' \quad (3.14)$$

1533 where  $\phi_u$  is a differentiable *update* function,  $\square_j$  is a differentiable *aggregation* function and  $\phi_m$  is a

1534 differentiable *message* function.  $\mathcal{N}'_i = \{n_j \in \mathcal{N} | (n_i, n_j) \in \mathcal{E}\}$  is the set of neighbours of  $n_i$ , i.e. the  
 1535 nodes  $n_j$  from which it exist an edge  $e_{i,j} \rightarrow (n_i, n_j)$ .  $k$  is the layer on which the message passing  
 1536 algorithm is applied. The update function need also a few other property if we want to keep the  
 1537 graph property, most notably the permutational invariance of its parameters (example: mean, std,  
 1538 sum, ...). The differents message, update and aggregation functions can really be any kind of function  
 1539 if they follow the constraint presented before, even small Neural Network.

1540 The edges features can also be updated, either by directly taking the results of  $\phi_m$  or by using another  
 1541 message function  $\phi_e$ .

1542 To explain this process, let's take the situation presented in Figure 3.13. We start with an input graph  
 1543 on left, in this case the message passing algorithm is mixing the color on each nodes and produce  
 1544 nodes of mixed color. For simplicity, the  $\phi_m$  and  $\phi_u$  function are the identity, they take a color and  
 1545 output the same color.

1546 Let's look at what's happening in the node 4. It has 3 neighbours and is a neighbour of itself. The four  
 1547 resulting  $\phi_m$  extract the color of each nodes and then feed them to the  $\square$  function. The  $\square$  function  
 1548 just equally distribute the color in the node. Finally the  $\phi_u$  function just update the node with the  
 1549 output of  $\square$ .

1550 Interestingly we see that the new node 4 does not have any yellow, the color of node 1. But if we were  
 1551 to run the message passing algorithm again, it would get some as node 2 is now partially yellow. If  
 1552 color here represent information, we see that multiple step are needed so that each node is "aware"  
 1553 of the informations the other nodes possess.

1554 Message passing is a very generic way of describing the process of GNN and it can be specialized for  
 1555 convolutional filtering [104], diffusion [105] and many other specific operation. GNN are used in a  
 1556 wide variety of application such as regression problematics, node classification, edge classification,  
 1557 node and edge prediction, ...

1558 It is a very versatile but complex tool.

### 1559 3.2.4 Adversarial Neural Network (ANN)

1560 The adversarial machine learning, Adversarial Neural Networks (ANN) in the case of neural net-  
 1561 work, is a family of unsupervised machine learning algorithms where the learning algorithm (gen-  
 1562 erator) is competing against another algorithm (discriminator). Taking the example of Generative  
 1563 Adversarial Networks, concept initially developed by Goodfellow et al. [106], the discriminator goal  
 1564 is to discriminate between data coming from a reference dataset and data produced by the generator.  
 1565 The generator goal, on the other hand, is to produce data that the discriminator would not be able to  
 1566 differentiate from data from the reference dataset. The expression of duality between the two models  
 1567 is represented in the loss where, at least a part of it, is driven by the results of the discriminator.

## 1568 3.3 State of the art of the Offline IBD reconstruction in JUNO

1569 The main reconstruction method currently run in JUNO is OMILREC, a data-driven method based  
 1570 on a likelihood maximization [107, 108] using only the LPMTs. The first step is to reconstruct the  
 1571 interaction vertex from which the energy reconstruction is dependent. It is also necessary for event  
 1572 pairing and classification.

1573 **3.3.1 Interaction vertex reconstruction**

1574 To start the likelihood maximization, a rough estimation of the vertex and of the event timing is  
 1575 needed. We start by estimating the vertex position using a charge based algorithm.

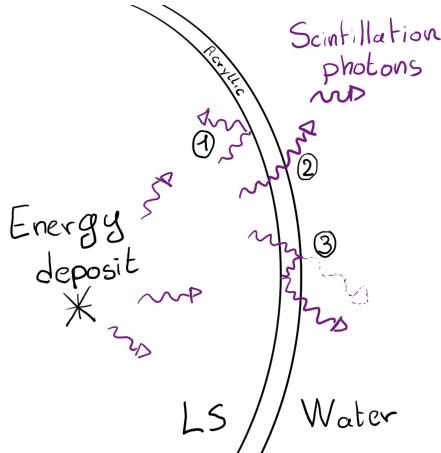
1576 **Charge based algorithm**

1577 The charge-based algorithm is basically base on the charge-weighted average of the PMT position.

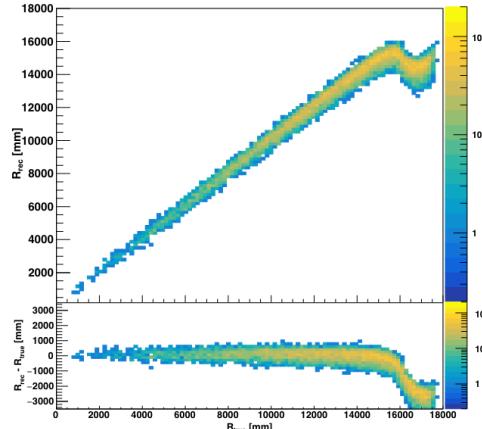
$$\vec{r}_{cb} = a \cdot \frac{\sum_i q_i \cdot \vec{r}_i}{\sum_i q_i} \quad (3.15)$$

1578 Where  $q_i$  is the reconstructed charge of the pulse of the  $i$ th PMT and  $\vec{r}_i$  is its position.  $\vec{r}_0$  is the  
 1579 reconstructed interaction position.  $a$  is a scale factor introduced because a weighted average over  
 1580 a 3D sphere is inherently biased. Using calibration we can estimate  $a \approx 1.3$  [109]. The results in  
 1581 Figure 3.14b shows that the reconstruction is biased from around 15m and further. This is due to the  
 1582 phenomena called “total reflection area” or TR Area.

1583 As depicted in the Figure 3.14a the optical photons, given that they have a sufficiently large incidence  
 1584 angle, can be deviated of their trajectories when passing through the interfaces LS-acrylic and water-  
 1585 acrylic due to the optical index difference. This cause photons to be lost or to be detected by PMT  
 1586 further than anticipated if we consider their rectilinear trajectories. This cause the charge barycenter  
 1587 the be located closer to the center than the event really is.



(A) Illustration of the different optical photons reflection scenarios. 1 is the reflection of the photon at the interface LS-acrylic or acrylic-water. 2 is the transmission of the photons through the interfaces. 3 is the conduction of the photon in the acrylic.



(B) Heatmap of  $R_{rec}$  and  $R_{rec} - R_{true}$  as a function of  $R_{true}$  for 4MeV prompt signals uniformly distributed in the detector calculated by the charge based algorithm

FIGURE 3.14

1588 It is to be noted that charge based algorithm, in addition to be biased near the edge of the detector,  
 1589 does not provide any information about the timing of the event. Therefore, a time based algorithm  
 1590 needs to be introduced to provide initial values.

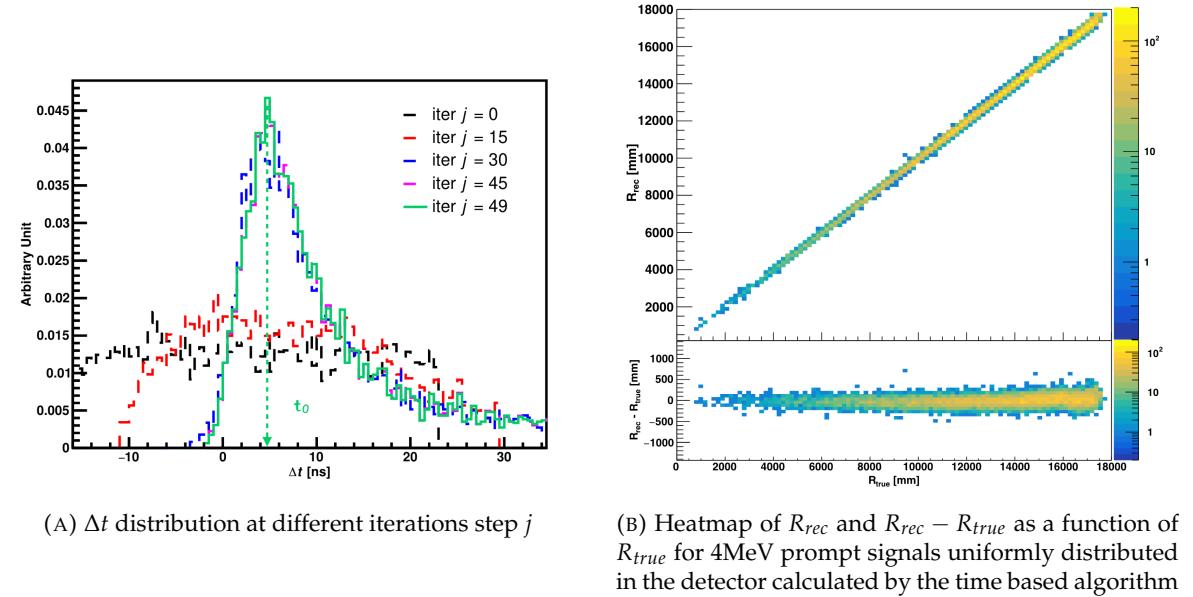


FIGURE 3.15

### 1591 Time based algorithm

1592 The time based algorithm use the distribution of the time of flight corrections  $\Delta t$  (Eq 3.16) of an event  
 1593 to reconstruct its vertex and  $t_0$ . It follow the following iterations:

- 1594 1. Use the charge based algorithm to get an initial vertex to start the iteration.  
 1595 2. Calculate the time of flight correction for the  $i$ th PMT using

$$\Delta t_i(j) = t_i - \text{tof}_i(j) \quad (3.16)$$

1596 where  $j$  is the iteration step,  $t_i$  is the timing of the  $i$ th PMT, and  $\text{tof}_i$  is the time-of-flight of the  
 1597 photon considering an rectilinear trajectory and an effective velocity in the LS and water (see  
 1598 [109] for detailed description of this effective velocity). Plot the  $\Delta t$  distribution and label the  
 1599 peak position as  $\Delta t^{\text{peak}}$  (see fig 3.15a).

- 1600 3. Calculate a correction vector  $\vec{\delta}[\vec{r}(j)]$  as

$$\vec{\delta}[\vec{r}(j)] = \frac{\sum_i \left( \frac{\Delta t_i(j) - \Delta t^{\text{peak}}(j)}{\text{tof}_i(j)} \right) \cdot (\vec{r}_0(j) - \vec{r}_i)}{N^{\text{peak}}(j)} \quad (3.17)$$

1601 where  $\vec{r}_0$  is the vertex position at the beginning of this iteration,  $\vec{r}_i$  is the position of the  $i$ th PMT.  
 1602 To minimize the effect of scattering, dark noise and reflection, only the pulse happening in a  
 1603 time window (-10 ns, +5 ns) around  $\Delta t^{\text{peak}}$  are considered.  $N^{\text{peak}}$  is the number of PE collected  
 1604 in this time-window.

- 1605 4. if  $\vec{\delta}[\vec{r}(j)] < 1\text{mm}$  or  $j \geq 100$ , stop the iteration. Otherwise  $\vec{r}_0(j+1) = \vec{r}_0(j) + \vec{\delta}[\vec{r}(j)]$  and go to  
 1606 step 2.

1607 However because the earliest arrival time is used,  $t_i$  is related to the number photoelectrons  $N_i^{\text{pe}}$   
 1608 detected by the PMT [110–112]. To reduce bias in the vertex reconstruction, the following equation

<sup>1609</sup> is used to correct  $t_i$  into  $t'_i$ :

$$t'_i = t_i - p_0 / \sqrt{N_i^{\text{pe}}} - p_1 - p_2 / N_i^{\text{pe}} \quad (3.18)$$

<sup>1610</sup> The parameters  $(p_0, p_1, p_2)$  were optimized to  $(9.42, 0.74, -4.60)$  for Hamamatsu PMTs and  $(41.31,$   
<sup>1611</sup>  $-12.04, -20.02)$  for NNVT PMTs [109]. The results presented in Figure 3.15b shows that the time  
<sup>1612</sup> based algorithm provide a more accurate vertex and is unbiased even in the TR area. This results  
<sup>1613</sup>  $(\vec{r}_0, t_0)$  is used as initial value for the likelihood algorithm.

#### <sup>1614</sup> Time likelihood algorithm

<sup>1615</sup> The time likelihood algorithm use the residual time expressed as follow

$$t_{\text{res}}^i(\vec{r}_0, t_0) = t_i - \text{tof}_i - t_0 \quad (3.19)$$

<sup>1616</sup> In a first order approximation, the scintillator time response Probability Density Function (PDF) can  
<sup>1617</sup> be described as the emission time profile of the scintillation photons, the Time Transit Spread (TTS)  
<sup>1618</sup> and the dark noise of the PMTs. The emission time profile  $f(t_{\text{res}})$  is described like

$$f(t_{\text{res}}) = \sum_k \frac{\rho_k}{\tau_k} e^{-\frac{t_{\text{res}}}{\tau_k}}, \sum_k \rho_k = 1 \quad (3.20)$$

<sup>1619</sup> as the sum of the  $k$  component that emit light in the LS each one characterised by it's decay time  $\tau_k$   
<sup>1620</sup> and intensity fraction  $\rho_k$ . The TTS component is expressed as a gaussian convolution

$$g(t_{\text{res}}) = \frac{1}{\sqrt{2\pi}\sigma} e^{-\frac{(t_{\text{res}}-\nu)^2}{2\sigma^2}} \cdot f(t_{\text{res}}) \quad (3.21)$$

<sup>1621</sup> where  $\sigma$  is the TTS of PMTs and  $\nu$  is the average transit time. The dark noise is not correlated with any  
<sup>1622</sup> physical events and considered as constant rate over the time window considered  $T$ . By normalizing  
<sup>1623</sup> the dark noise probability  $\epsilon(t_{\text{res}})$  as  $\int_T \epsilon(t_{\text{res}}) dt_{\text{res}} = \epsilon_{\text{dn}}$ , it can be integrated in the PDF as

$$p(t_{\text{res}}) = (1 - \epsilon_{\text{dn}}) \cdot g(t_{\text{res}}) + \epsilon(t_{\text{res}}) \quad (3.22)$$

<sup>1624</sup> The distribution of the residual time  $t_{\text{res}}$  of an event can then be compared to  $p(t_{\text{res}})$  and the best  
<sup>1625</sup> fitting vertex  $\vec{r}_0$  and  $t_0$  can be chosen by minimizing

$$\mathcal{L}(\vec{r}_0, t_0) = -\ln \left( \prod_i p(t_{\text{res}}^i) \right) \quad (3.23)$$

<sup>1626</sup> The parameter of Eq. 3.22 can be measured experimentally. The results shown in Figure 3.16  
<sup>1627</sup> used PDF from monte carlo simulation. The results shows that  $R_{\text{rec}} - R_{\text{true}}$  is biased depending  
<sup>1628</sup> on the energy. While this could be corrected using calibration, another algorithm based on charge  
<sup>1629</sup> likelihood was developed to correct this problem.

#### <sup>1630</sup> Charge likelihood algorithm

<sup>1631</sup> Similarly to the time likelihood algorithms that use a timing PDF, the charge likelihood algorithm  
<sup>1632</sup> use a PE PDF for each PMT depending on the energy and position of the event. With  $\mu(\vec{r}_0, E)$  the  
<sup>1633</sup> mean expected number of PE detected by each PMT, the probability to observe  $N_{\text{pe}}$  in a PMT follow  
<sup>1634</sup> a Poisson distribution. Thus

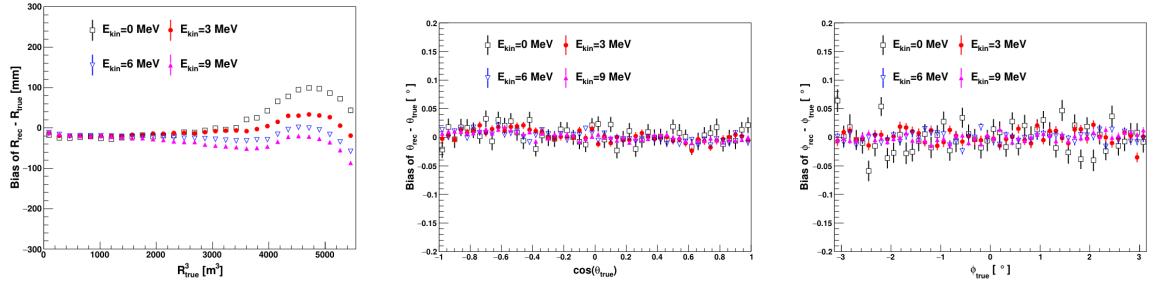


FIGURE 3.16 – Bias of the reconstructed radius  $R$  (left),  $\theta$  (middle) and  $\phi$  (right) for multiple energies by the time likelihood algorithm

- The probability to observe no hit ( $N_{pe} = 0$ ) in the  $j$ th PMT is  $P_{nohit}^j(\vec{r}_0, E) = e^{-\mu_j}$
- The probability to observe  $N_{pe} \neq 0$  in the  $i$ th PMT is  $P_{hit}^i(\vec{r}_0, E) = \frac{\mu^{N_{pe}} e^{-\mu_i}}{N_{pe}^i!}$

Therefore, the probability to observe a specific hit pattern can be expressed as

$$P(\vec{r}_0, E) = \prod_j P_{nohit}^j(\vec{r}_0, E) \cdot \prod_i P_{hit}^i(\vec{r}_0, E) \quad (3.24)$$

The best fit values of  $\vec{R}_0$  and  $E$  can then be calculated by minimizing the negative log-likelihood

$$\mathcal{L}(\vec{r}_0, E) = -\ln(P(\vec{r}_0, E)) \quad (3.25)$$

In principle,  $\mu_i(\vec{r}_0, E)$  could be expressed

$$\mu_i(\vec{r}_0, E) = Y \cdot \frac{\Omega(\vec{r}_0, r_i)}{4\pi} \cdot \epsilon_i \cdot f(\theta_i) \cdot e^{-\sum_m \frac{d_m}{\zeta_m}} \cdot E + \delta_i \quad (3.26)$$

where  $Y$  is the energy scale factor,  $\Omega(\vec{r}_0, r_i)$  is the solid angle of the  $i$ th PMT,  $\epsilon_i$  is its detection efficiency,  $f(\theta_i)$  its angular response,  $\zeta_m$  is the attenuation length in the materials and  $\delta_i$  the expected number of dark noise.

However Eq. 3.26 assume that the scintillation light yield is linear with energy and describe poorly the contribution of indirect light, shadow effect due to the supporting structure and the total reflection effects. The solution is to use data driven methods to produce the pdf by using the calibrations sources and position described in Section 2.4. In the results presented in Figures 3.17, the PDF was produced using MC simulation and 29 specific calibrations position [109] along the Z-axis of the detector. We see that the charge likelihood algorithm show little bias in the TR area and a better resolution than the time likelihood. The Figure 3.18 shows the radial resolution of the different algorithm presented for this section, we can see the refinement at each step and that the charge likelihood yield the best results.

The charge based likelihood algorithms already give some information on the energy as Eq. 3.25 is minimized but the energy can be further refined as shown in the next section.

### 3.3.2 Energy reconstruction

As explained in Section 2.1.1, energy resolution is crucial for the NMO and oscillation parameters measurements. Thus the energy reconstruction algorithm should take into consideration as much

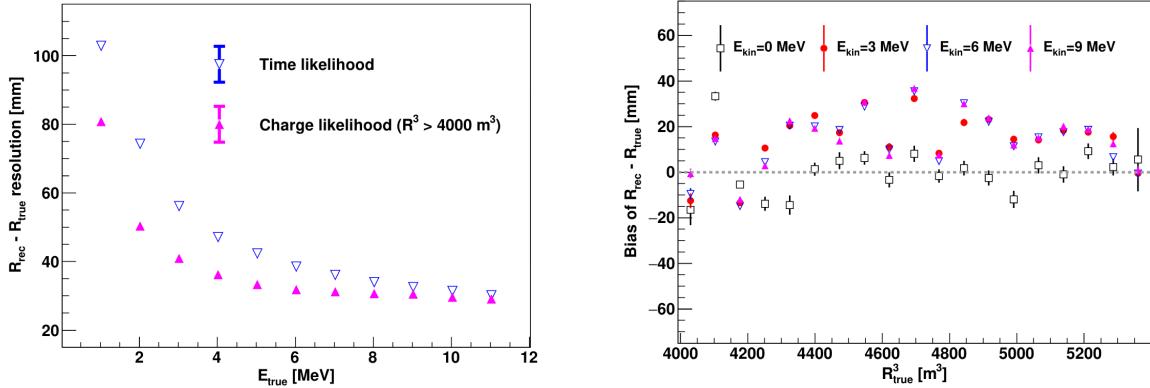


FIGURE 3.17 – On the left: Resolution of the reconstructed  $R$  as a function of the energy in the TR area ( $R^3 > 4000 \text{ m}^3 \equiv R > 16 \text{ m}$ ) by the charge and time likelihood algorithms. On the right: Bias of the reconstructed  $R$  in the TR area for different energies by the charge likelihood algorithm

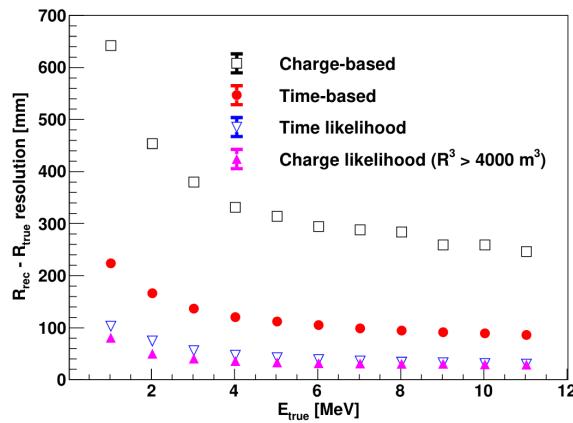


FIGURE 3.18 – Radial resolution of the different vertex reconstruction algorithms as a function of the energy

1657 detector effect as possible. The following method is a data driven method based on calibration  
1658 samples inspired by the charge likelihood algorithm described above [113].

### 1659 Charge estimation

1660 The most important element in the energy reconstruction is  $\mu_i(\vec{r}_0, E)$  described in Eq. 3.26. For  
1661 realistic cases, we also need to take into account the electronics effect that were omitted in the  
1662 previous section. Those effect will cause a charge smearing due to the uncertainties in the  $N_{pe}$   
1663 reconstruction. Thus we define  $\hat{\mu}^L(\vec{r}_0, E)$  which is the expected  $N_{pe}/E$  in the whole detector for an  
1664 event with visible energy  $E_{vis}$  and position  $\vec{r}_0$ . The position of the event and PMTs are now defined  
1665 using  $(r, \theta, \theta_{pmt})$  as defined in Figure 3.19b.

$$\hat{\mu}(r, \theta, \theta_{pmt}, E_{vis}) = \frac{1}{E_{vis}} \frac{1}{M} \sum_i^M \frac{q_i}{Q_i} - \mu_i^D, \quad \mu_i^D = \text{DNR}_i \cdot L \quad (3.27)$$

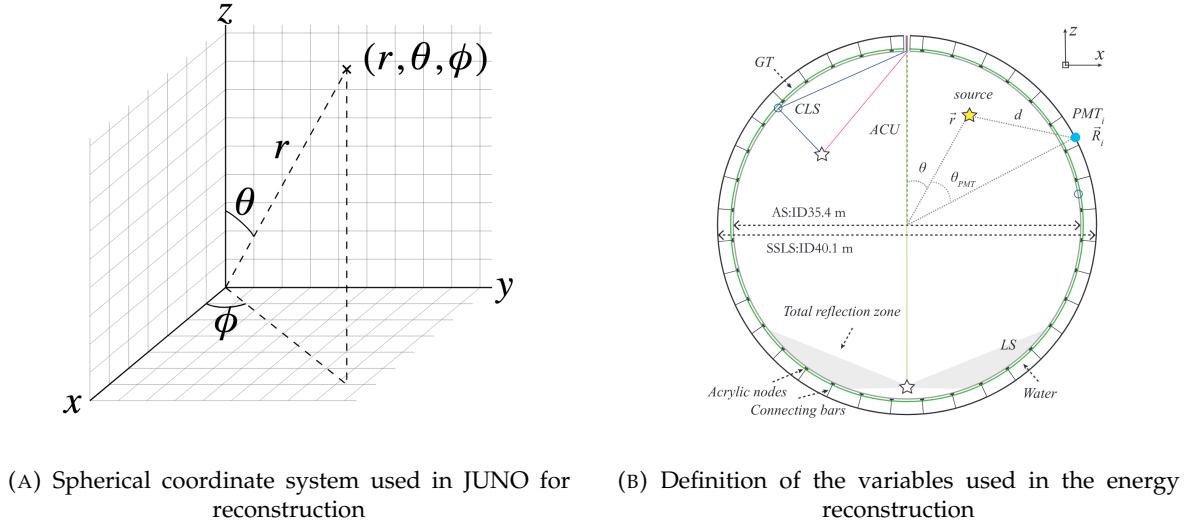


FIGURE 3.19

where  $i$  runs over the PMTs with the same  $\theta_{pmt}$ ,  $DE_i$  is the detection efficiency of the  $i$ th PMT.  $\mu_i^D$  is the expected number of dark noise photoelectrons in the time window  $L$ . The time window have been optimized to  $L = 280$  ns [113].  $\bar{q}_i$  is the average recorded photoelectrons in the time window and  $\hat{Q}_i$  is the expected average charge for 1 photoelectron. The  $N_{pe}$  map is constructed following the procedure described in [108].

### Time estimation

The second important observable is the hit time of photons that was previously defined in Eq. 3.19. It is here refined as

$$t_r = t_h - \text{tof} - t_0 = t_{LS} + t_{TT} \quad (3.28)$$

where  $t_h$  is the time of hit,  $t_{LS}$  is the scintillation time and  $t_{TT}$  the transit time of PMTs that is described by a gaussian

$$t_{TT} = \mathcal{N}(\overline{\mu_{TT} + t_d}, \sigma_{TT}) \quad (3.29)$$

where  $\mu_{TT}$  is the mean transit time in PMTs,  $\sigma_{TT}$  is the Transit Time Spread (TTS) of the PMTs and  $t_d$  is the delay time in the electronics. The effective refraction index of the LS is also corrected to take into account the propagation distance in the detector.

The timing PDF  $P_T(t_r|r, d, \mu_l, \mu_d, k)$  can now be generated using calibration sources [113]. This PDF describe the probability that the residual time of the first photon hit is in  $[t_r, t_r + \delta]$  with  $r$  the radius of the event vertex,  $d = |\vec{r} - \vec{r}_{PMT}|$  the propagation distance,  $\mu_l$  and  $\mu_d$  the expected number of PE and dark noise in the electronic reading window and  $k$  is the detected number of PE.

Now let denote  $f(t, r, d)$  the probability density function of "photoelectron hit a time t" for an event happening at  $r$  where the photons traveled the distance  $d$  in the LS

$$F(t, r, d) = \int_t^L f(t', r, d) dt' \quad (3.30)$$

Based on the PDF for one photon  $k = 1$ , one can define

$$P_T^l(t|k = n) = I_n^l[f_l(t)F_l^{n-1}(t)] \quad (3.31)$$

1686 where the indicator  $l$  means that the photons comes from the LS and  $I_n^l$  a normalisation factor. To this  
 1687 pdf we add the probability to have photons coming from the dark noise indicated by the indicator  $d$   
 1688 using

$$f_d(t) = 1/L, F_d(t) = 1 - \frac{t}{L} \quad (3.32)$$

1689 and so for the case where only one photon is detected by the PMT ( $k = 1$ )

$$P_T(t|\mu_l, \mu_d, k=1) = I_1[P(1, \mu_l)P(0, \mu_d)f_l(t) + P(0, \mu_l)P(1, \mu_d)f_d(t)] \quad (3.33)$$

1690 where  $P(k_\alpha, \mu_\alpha)$  is the Poisson probability to detect  $k_\alpha$  PE from  $\alpha \in \{l, d\}$  with the condition  $k_l + k_d =$   
 1691  $k$ .

1692 Now that we have the individual timing and charge probability we can construct the charge likeli-  
 1693 hood referred as QMLE:

$$\mathcal{L}(q_1, q_2, \dots, q_N | \vec{r}, E_{vis}) = \prod_{j \in \text{unfired}} e^{-\mu_j} \prod_{i \in \text{fired}} \left( \sum_{k=1} P_Q(q_i|k) \cdot P(k, \mu_i) \right) \quad (3.34)$$

1694 where  $\mu_i = E_{vis}\hat{\mu}_i^L + \mu_i^D$  and  $P(k, \mu_i)$  is the Poisson probability of observing  $k$  PE.  $P_Q(q_i|k)$  is the  
 1695 charge pdf for  $k$  PE. And we can also construct the time likelihood referred as TMLE:

$$\mathcal{L}(t_{1,r}, t_{2,r}, \dots, t_{N,r} | \vec{r}, t_0) = \prod_{i \in \text{hit}} \frac{\sum_{k=1}^K P_T(t_{i,r}|r, d, \mu_i^l, \mu_i^d, k) \cdot P(k, \mu_i^l + \mu_i^d)}{\sum_{k=1}^K P(k, \mu_i^l + \mu_i^d)} \quad (3.35)$$

1696 where  $K$  is cut to 20 PE and hit is the set of hits satisfying  $-100 < t_{i,r} < 500$  ns.

1697 Merging those two likelihood give the charge-time likelihood QTMLE, the core algorithm of OMIL-  
 1698 REC.

$$\mathcal{L}(q_1, q_2, \dots, q_N; t_{1,r}, t_{2,r}, \dots, t_{N,r} | \vec{r}, t_0, E_{vis}) = \mathcal{L}(q_1, q_2, \dots, q_N | \vec{r}, E_{vis}) \cdot \mathcal{L}(t_{1,r}, t_{2,r}, \dots, t_{N,r} | \vec{r}, t_0) \quad (3.36)$$

1699 The radial and energy resolutions of the different likelihood are presented in Figure 3.20 (from [113]).  
 1700 We can see the improvement of adding the time information to the vertex reconstruction and that  
 1701 an increase in vertex precision can bring improvement in the energy resolution, especially at low  
 1702 energies.

1703 Data driven methods prove to be performant in the energy and vertex reconstruction given that we  
 1704 have enough calibrations sources to produce the PDF. In addition to this, member of JUNO have  
 1705 developed ML algorithms for reconstruction. The one focused on IBD reconstruction are presented  
 1706 in the next section.

### 1707 3.3.3 Machine learning for reconstruction

1708 The power of ML is the ability to model complex response to a specific problem. In JUNO the  
 1709 reconstruction problematic can be expressed as follow: knowing that each PMT, large or small,  
 1710 detected a given number of PE  $Q$  at a given time  $t$  and their position is  $x, y, z$  where did the energy  
 1711 was deposited and how much energy was it, modeling a function that naively goes:

$$\mathbb{R}^{5 \times N_{pmt}} \mapsto \mathbb{R}^4 \quad (3.37)$$

1712 It is worth pointing that while this is already a lot in informations, this is not the rawest representa-  
 1713 tion of the experiment. We could indeed replace the charge and time by the waveform in the time  
 1714 window of the event but that would lead to an input representation size that would exceed our  
 1715 computational limits. Also, due to those computational limits, most of the ML algorithm reduce this

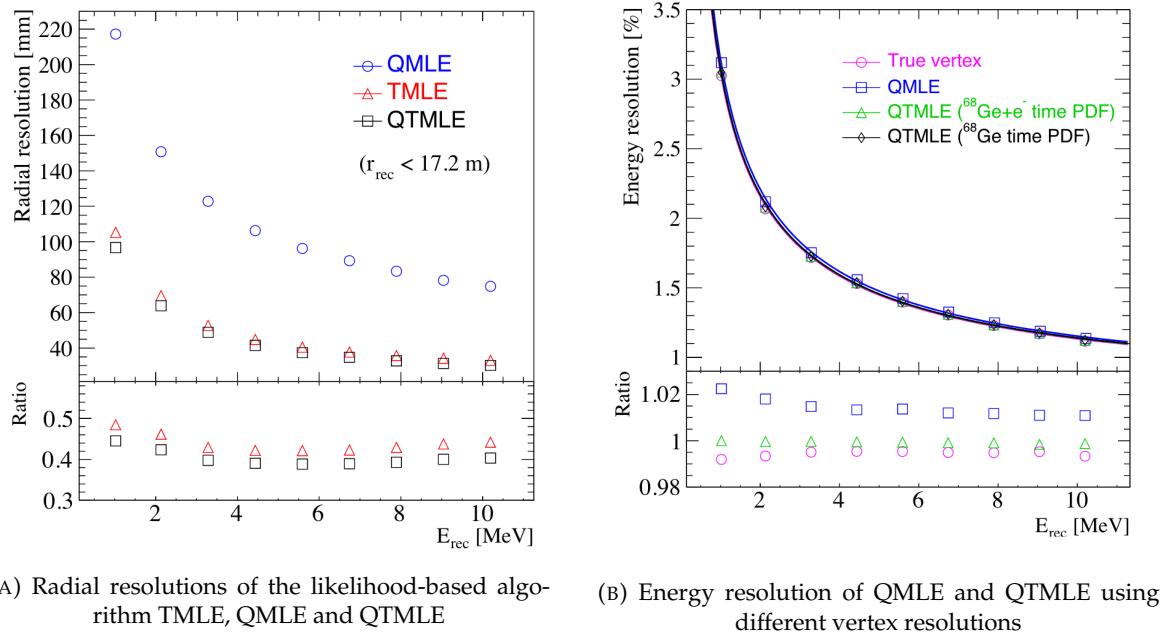


FIGURE 3.20

1716 input phase space either by structurally encoding the information (pictures, graph), by aggregating  
 1717 it (mean, variance, ...) or by exploiting invariance and equivariance of the experiment (rotational  
 1718 invariance due to the sphericity, ...).

1719 For machine learning to converge to performant algorithm, a large dataset exploring all the phase  
 1720 space of interest is needed. For the following studies, data from the monte carlo simulation presented  
 1721 in Section 2.6 are used for training. When the detector will be finished calibrations sources will be  
 1722 complementarily be used.

### 1723 Boosted Decision Tree (BDT)

1724 On of the most classic ML method used in physics in last years is the Boosted Decision Tree. They  
 1725 have been explored for vertex reconstruction [114] et for energy reconstruction [114, 115].

1726 For vertex and energy reconstruction a BDT was developed using the aggregated informations pre-  
 1727 sented in 3.1.

Parameter	description
$n_{\text{Hits}}$	Total number of hits
$x_{cc}, y_{cc}, z_{cc}, R_{cc}$	Coordinates of the center of charge
$ht_{\text{mean}}, ht_{\text{std}}$	Hit time mean and standard deviation

TABLE 3.1 – Features used by the BDT for vertex reconstruction

1728 Its reconstruction performances are presented in Figure 3.22.

1729 A second and more advanced BDT, subsequently named BDTE, that only reconstruct energy use a  
 1730 different set of features [115]. They are presented in the table 3.2

AccumCharge	$ht_{5\%-2\%}$
$R_{cht}$	$pe_{mean}$
$z_{cc}$	$J_{cht}$
$pe_{std}$	$\phi_{cc}$
nPMTs	$ht_{35\%-30\%}$
$ht_{kurtosis}$	$ht_{20\%-15\%}$
$ht_{25\%-20\%}$	$pe_{35\%}$
$R_{cc}$	$ht_{30\%-25\%}$

TABLE 3.2 – Features used by the BDTE algorithm.  $pe$  and  $ht$  reference the charge and hit-time distribution respectively and the percentages are the quantiles of those distributions.  $cht$  and  $cc$  reference the barycenters of hit time and charge respectively

### 1731 Neural Network (NN)

1732 Three type of neural networks have explored for event reconstruction in JUNO Deep Neural Network (DNN), Convolutional Neural Network (CNN) and Graph Network (GNN).  
 1733

1734 The CNN are using 2D projection of the detector representing it as an image with two channel, one  
 1735 for the charge  $Q$  and one for the time  $t$ . The position of the PMTs is structurally encoded in the pixel  
 1736 containing the information of this PMT. In [114], the pixel is chosen based on a transformation of  $\theta$   
 1737 and  $\phi$  coordinates to the 2D plane and rounded to the nearest pixel. A sufficiently large image has  
 1738 been chosen to prevent two PMT to be located in the same pixel. An example of this projection can  
 1739 be found in Figure 3.21. The performances of the CNN can be found in Figure 3.22.

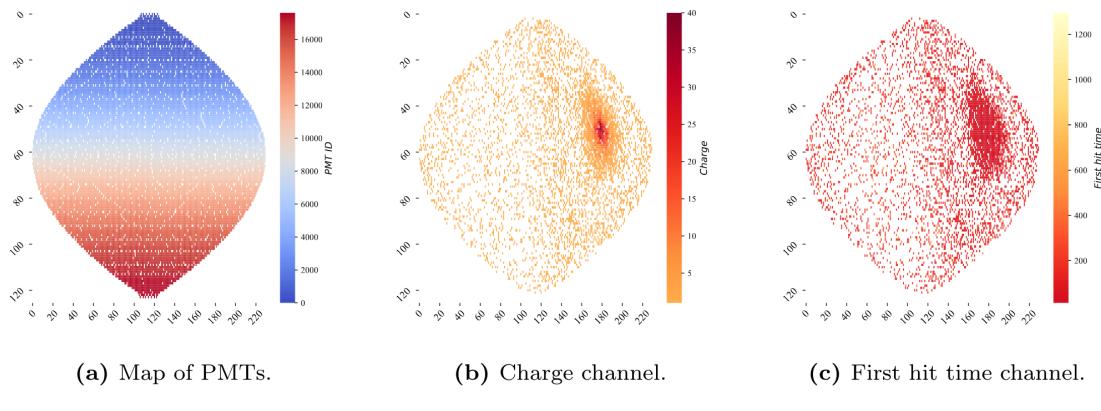


FIGURE 3.21 – Projection of the LPMTs in JUNO on a 2D plane. (a) Show the distribution of all PMTs and (b) and (c) are example of what the charge and time channel looks like respectively

1740 Using 2D have the upside of encoding a large part of the informations structurally but loose the  
 1741 rotational invariance of the detector. It also give undefined information to the neural network  
 1742 (what is a pixel without PMT ? What should be its charge and time ?), cause deformation in the  
 1743 representation of the detector (sides of projection) and loose topological informations.  
 1744

One of the way to present structurally the sphericity of JUNO to a NN is to use a graph: A collection  
 1745 of objects  $V$  called nodes and relations  $E$  called edges, each relation associated to a couple  $v_1, v_2$   
 1746 forming the graph  $G(E, V)$ . Nodes and edges can hold informations or features. In [114] the nodes,  
 1747 are geometrical region of the detector as defined by the HealPix [116]. The features of the nodes are  
 1748 aggregated informations from the PMTs it contains. The edges contains geographic informations of  
 1749 the nodes relative positions.

1750 This data representation has the advantages to keep the topology of the detector intact. It also permit  
 1751 the use of rotational invariant algorithms for the NN, thus taking advantage of the symmetries of the  
 1752 detector.

1753 The neural network then process the graph using Chebyshev Convolutions [104]. The performances  
 1754 of the GNN are presented in Figure 3.22.

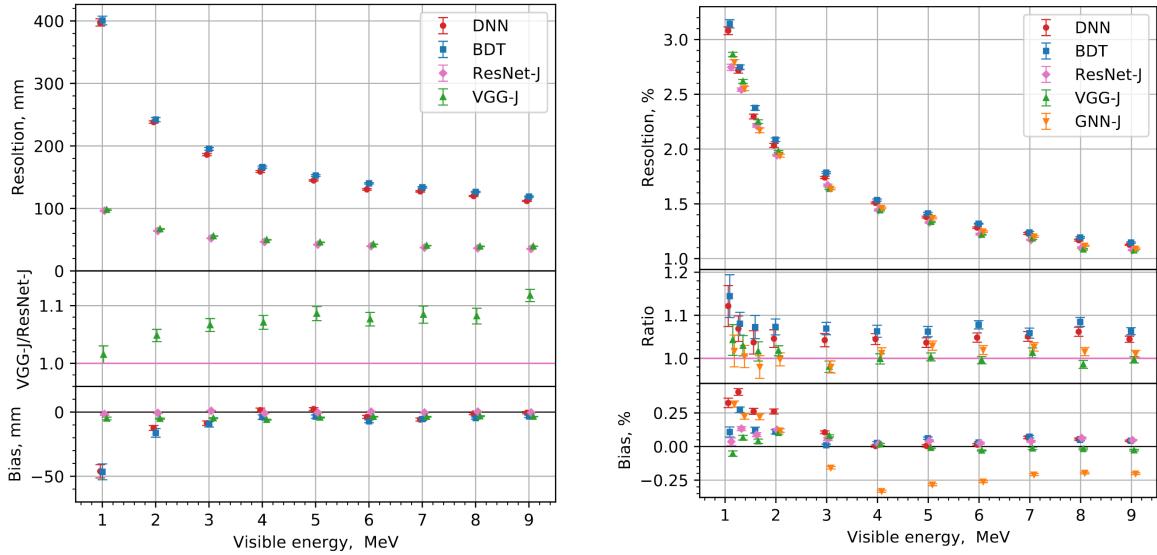


FIGURE 3.22 – Radial (left) and energy (right) resolutions of different ML algorithms.  
 The results presented here are from [114]. DNN is a deep neural network, BDT is a  
 BDT, ResNet-J and VGG-J are CNN and GNN-J is a GNN.

1755 Overall ML algorithms show similar performances as classical algorithms in term of energy recon-  
 1756 structions with the more complex structure CNN and GNN showing better performances than BDT  
 1757 and DNN. For vertex reconstruction, the BDT and DNN show poor performance while CNN are on  
 1758 the level of the classical algorithms.

## 1759 3.4 Conclusion

1760 That these first DL algorithms tried at JUNO to reconstruct IBDs do not outperform the classical  
 1761 method can be explained. They constitute a first exploration of these methods potential, as do the  
 1762 original GNN we describe in Chapter 5. Indeed, the likelihood method is also based on the full list of  
 1763 the charges ( $Q$ ) and times ( $t$ ) all PMTs, and the PDF's design accounts for an advanced knowledge  
 1764 of the detector (with a lot of human expertise). The fact that the methods presented in this chapter  
 1765 can learn enough from just the  $Q, t$  list, to reach similar performance, is already an interesting result.  
 1766 But this is not decisive yet, in my opinion.

1767 Actually, is there hope that one day DL methods reach better results at JUNO than classical's ? This  
 1768 is not a trivial question. A possibility would be to let them start from an even rawer level (involving a  
 1769 number of variables which would make a likelihood intractable). This would mean, instead of  $Q$  and  
 1770  $t$ , the full waveform in each PMT. With such a quantity of input information to analyse to identify  
 1771 patterns, even DL methods can be limited. The choice of architecture is then important, to guide the  
 1772 algorithm towards pertinent features. We doubt whether CNN's would be the best choice here. We  
 1773 bet that GNN's could be better tools, with more flexibility to hierachise information (the choice of  
 1774 which PMTs to link already helps here, as well as the possible usage of higher order quantities). The

1775 first GNN developped in JUNO (described above, [114]) does not do that. It's still only based on  $(Q, t)$   
1776 couples and link only neighbour PMTs in its first layer. It serves essentially as a way to avoid the  
1777 problems encountered by CNNs due to the planar projection of a spherical image.

1778 In chapter 5, we tried an original GNN architecture. The goal was not yet to include a rawer  
1779 information, but to see if this architecture would perform as well as the one described above when  
1780 using  $Q$ 's and  $t$ 's as the rawest information. If so, then there is hope that when rawer information  
1781 will be included, this orginal architecture will be the one able to best use it.

<sup>1782</sup> **Chapter 4**

<sup>1783</sup> **Image recognition for IBD  
reconstruction with the SPMT system**

*Dave - Give me the position and momentum, HAL.*

*HAL - I'm afraid I can't do that Dave.*

*Dave - What's the problem ?*

*HAL - I think you know what the problem is just as well as I do.*

*Dave - What are you talking about, HAL?*

*HAL -  $\sigma_x \sigma_p \geq \frac{\hbar}{2}$*

<sup>1786</sup> **Contents**

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<sup>1800</sup> As explained in Chapter 2, JUNO is an experiment composed of two systems, the Large Photomultiplier (LPMT) system and the Small Photomultiplier (SPMT) system. Both of them observe the same physics events inside of the same medium but they differ in their photo-coverage, respectively 75.2% and 2.7%, their dynamic range (see Section 2.3.2), a thousands versus a few dozen, and their front-end electronics (see section 2.3.2).

<sup>1801</sup> The SPMT system is essential to the deployment of the Dual Calorimetry techniques, already mentioned in Section 3.3 and described in [82, 84, 117]. It is indeed less subject than the LPMTs to charge non linearity effects (QNL). This topic will be studied in more detail in Chapter 7, where the potential of one of the Dual Calorimetry techniques is explored. It consists on combined oscillation analyses based on two antineutrino energy spectra : one reconstructed with the LPMT system, the other one with the SPMT system. For that purpose, it is therefore necessary to have reconstruction tools available. Well maintained tools using the LPMT are available in the collaboration's official software. This is not the case concerning the SPMT system, where algorithms were developed more sporadically. This is one of the reasons why we developed the CNN described in this chapter.

1816 Our efforts on it were limited to the early months of this thesis: it was above all a way to learn about  
 1817 ML and about JUNO's detector and software. We benchmarked its performance against a classical  
 1818 algorithm developed in Chapter 4 of [34] but not yet implemented in JUNO's software.

1819 As discussed in Chapter 3, Machine Learning (ML) algorithms shine when modeling highly dimen-  
 1820 sional data from a given dataset. In our case, we have access to complete monte-carlo simulation of  
 1821 our detector to produce large datasets that could represent multiple years of data taking. Ideally ML  
 1822 algorithms would be able to consider the entirety of the information in the detector and converge on  
 1823 the best parameters to yield optimal results.

1824 The difference between this ideal and what can be achieved in reality is an important subject. In  
 1825 particular, we wonder if an exhaustive usage of the information present in the detector could lead to  
 1826 use informations that are mismodelled in our simulated training samples (or present only in these  
 1827 samples) and therefore lead to biases when the algorithm is applied to real data. A simple way  
 1828 to start addressing this reliability issue is to try to evaluate to which extent various reconstruction  
 1829 methods use the same information. An attempt at this is presented at the end of this chapter. This is  
 1830 also the subject of Chapter 6.

## 1831 4.1 Method and model

1832 One of simplest way to look at JUNO data is to consider the detector as an array of geometrically  
 1833 distributed sensors on a sphere. Their repartition is almost homogeneous, on this sphere surface  
 1834 providing an almost equal amount of information per unit surface. It is then tempting to represent  
 1835 the detector as a spherical image with the PMTs in place of pixels. Two events with two different  
 1836 energy or position would produce two different images.

1837 The most common approach in machine learning for image processing and image recognition is the  
 1838 Convolutional Neural Network (CNN). It is widely used in research and industry [98, 118–120] due  
 1839 to its strengths (see Section 3.2.2) and has proven its relevance in image processing.

1840 Some CNN are developed to process spherical images [121] but for the sake of simplicity and as a  
 1841 first approach we decided to go with a planar projection of the detector, approach that has proven  
 1842 its efficiency using the LPMT system (see Section 3.3.3). The details about this planar projection will  
 1843 be discussed in section 4.1.2.

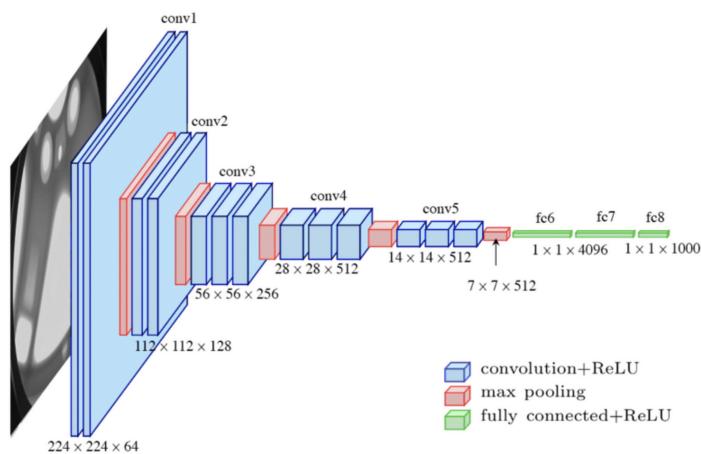


FIGURE 4.1 – Graphic representation of the VGG-16 architecture, presenting the different kind of layer composing the architecture.

### 4.1.1 Model

The architecture we use is derived from the VGG-16 architecture [98] illustrated in Figure 4.1. We define a set of hyperparameters that will define the size, complexity and computational power of the NN. The chose hyperparameters are detailed below and their values are presented in table 4.1.

- **N<sub>blocks</sub>**: the number of convolution blocks, a block being composed of two convolutional layers with  $3 \times 3$  filters using ReLU activation function, a  $3 \times 3$  kernel max-pooling layer (except for the last block).
- **N<sub>channels</sub>**: The number of channels in the first block. The number of channels in the subsequent blocks is computed using  $N_{channels}^i = i * N_{channels}$ ,  $i \in [1..N_{blocks}]$ .
- **FCDNN configuration**: The result of the last convolution layer is flattened then fed to a FCDNN. Its configuration is expressed as the ouputs of sequenced fully connected linear layer using the PReLU activation function. For example  $2 * 1024 + 2 * 512$  is the sequence of 2 layers which output is 1024 followed by 2 other layers with an output of 512. Finally the last layer is a linear layer outputing 4 features wihtout activation function. Each feature of the last layer represent a component of the interaction vertex: Energy, X, Y, Z.
- **Loss**: The loss function. In this work we study two different loss function ( $E + V$ ) and ( $E_r + V_r$ ) detailed below.

$$(E + V)(E, x, y, z) = (E - E_{dep})^2 + 0.85 \sum_{\lambda \in [x, y, z]} (\lambda - \lambda_{true})^2 \quad (4.1)$$

$$(E_r + V_r)(E, x, y, z) = \frac{(E - E_{dep})^2}{E_{dep}} + \frac{10}{R} \sum_{\lambda \in [x, y, z]} (\lambda - \lambda_{true})^2 \quad (4.2)$$

where  $E_{dep}$  is the deposited energy and  $R$  is the radius of JUNO's CD. With the energy in MeV and the distance in meters, we use the factor 0.85 and 10 to balance the two term of the loss function so they have the same magnitude.

The loss function ( $E + V$ ) is close to a simple Mean Squared Error (MSE). MSE is one of the most basic loss function, the derivative is simple and continuous in every point. It is a strong starting point to explore the possibility of CNNs. The loss ( $E_r + V_r$ ) can be seen as a relative MSE.

The idea is that: due to the inherent statistic uncertainty over the number of collected Number of Photo Electrons (NPE), the absolute resolution  $\sigma(E - E_{true})$  will be larger at higher energy than at low energy. But we expect the *relative* energy resolution  $\frac{\sigma(E - E_{true})}{E_{true}}$  to be smaller at high energy than lower energy as illustrated in Figure 3.20. Because of this, by using simple MSE the most important part in the loss come from the high energy part of the dataset whereas with a relative MSE, the most important part become the low energy events in the dataset. We hope that by using a relative MSE, the neural network will focus on low energy events where the reconstruction is considered the hardest.

The above losses and their parameters values results from fine-tuning after multiples runs and adjustments of the full random search.

Each combinations of those hyperparameters (for example ( $N_{blocks} = 2, N_{channels} = 32$ , FCDNN =  $(2 * 1024)$ , Loss =  $(E + V)$ )) produce models, hereinafter referred as configurations, are then tested and compared to each other over an analysis sample.

On top those generated models, we define 4 hand tailored models:

- Gen<sub>0</sub>:  $N_{blocks} = 4$ ,  $N_{channels} = 64$ , FCDNN configuration:  $1024 * 2 + 512 * 2$ , Loss  $\equiv E + V$
- Gen<sub>1</sub>:  $N_{blocks} = 4$ ,  $N_{channels} = 64$ , FCDNN configuration:  $1024 * 2 + 512 * 2$ , Loss  $\equiv E_r + V_r$
- Gen<sub>2</sub>:  $N_{blocks} = 5$ ,  $N_{channels} = 64$ , FCDNN configuration:  $4096 * 2 + 1024 * 2$ , Loss  $\equiv E + V$

- 1884 — Gen<sub>3</sub>:  $N_{blocks} = 5$ ,  $N_{channels} = 64$ , FCDNN configuration:  $4096 * 2 + 1024 * 2$ , Loss  $\equiv E_r + V_r$

1885  
 1886 The resulting models possess between 2'041'034, for Gen<sub>52</sub> and Gen<sub>53</sub>, and 5'759'839'242 parameters,  
 1887 for Gen<sub>26</sub> and Gen<sub>27</sub>. The models of interest in this thesis, from which the results are discussed  
 1888 in Section 4.3, possess 86'197'196 parameters for Gen<sub>30</sub> and 332'187'530 parameters for Gen<sub>42</sub>. For  
 1889 comparison the model of CNN developed in JUNO before posses 38'352'403 parameters [114].

$N_{blocks}$	{2, 3, 4}
$N_{channels}$	{32, 64, 128}
FCDNN configurations	$2 * 1024$ $2 * 2048 + 2 * 1024$ $3 * 2048 + 3 * 512$ $2 * 4096$
Loss	{ $E + V$ , $E_r + V_r$ }

TABLE 4.1 – Sets of hyperparameters values considered in this study

1890 To rank the various configuration we cannot used directly the mean loss over the validation dataset  
 1891 as  $(E + V)$  and  $(E_r + V_r)$  are not numerically comparable. We thus use the following quantities,  
 1892 directly related to the reconstruction performances:

- 1893 — The mean absolute energy error  $\langle E \rangle = \langle |E - E_{true}| \rangle$ . It is an indicator of the energy bias of our  
 1894 reconstruction.
- 1895 — The standard deviation of the energy error  $\sigma E = \sigma(E - E_{true})$ . This the indicator on our  
 1896 precision in energy reconstruction.
- 1897 — The mean distance between the reconstructed vertex and the true vertex  $\langle V \rangle = \langle |\vec{V} - \vec{V}_{true}| \rangle$ .  
 1898 This an indicator of the bias and precision of our vertex reconstruction.
- 1899 — The standard deviation of the distance between the true and reconstructed vertex  $\sigma V = \sigma |\vec{V} - \vec{V}_{true}|$ . This is an indicator if the precision in our vertex reconstruction.

1900  
 1901  
 1902 The models were developped in Python using the Pytorch framework [100] using NVIDIA A100  
 1903 [122] and NVIDIA V100 [123] gpus. The A100 was split in two, thus the accessible gpu memory  
 1904 was the same as V100, 20 Gb, making it impossible to train some of the architectures due to memory  
 1905 consumption.

1906 The training was monitored in realtime by a custom tooling that was developed during this thesis,  
 1907 DataMo [124].

1908 The training of one model takes between 4h and 15h depending of its size, overall training the full  
 1909 72 models takes around 500 GPU hours. Even with parallel training, this random search hyper-  
 1910 optimisation was time consuming.

### 1911 4.1.2 Data representation

1912 This data is represented as  $240 \times 240$  images with a charge  $Q$  channel and a time  $t$  channel. The  
 1913 SPMTs are then projected on the plane as illustrated in Figure 4.2b using the coordinate system  
 1914 presented in 4.2a. The  $P_y$  coordinate, the row corresponding to the SPMT in the projection, is  
 1915 proportional to  $\theta$ . The  $P_x$  coordinate, the column corresponding to the SPMT in the projection, is  
 1916 defined by  $\phi \sin \theta$  in spherical coordinates.  $\theta = 0$  is defined as being the top of the detector and  $\phi = 0$   
 1917 is defined as an arbitrary direction in the detector. In practice,  $\phi = 0$  is given by the MC simulation.

$$P_y = \left\lfloor \frac{\theta \cdot H}{\pi} \right\rfloor, \theta \in [0, \pi] \quad (4.3)$$

$$P_x = \left\lfloor \frac{(\phi + \pi) \sin \theta \cdot W}{2\pi} \right\rfloor, \phi \in [-\pi, \pi], \theta \in [0, \pi] \quad (4.4)$$

where  $H$  is the height of the image,  $W$  the width of the image and  $(0, 0)$  the top left corner of the image.

This projection keep the SPMT position in the image proportional to their spherical coordinates while keeping the neighbouring information. This proportionality allow us to keep the specificities of the detector structure, the vertical bands visible in 4.2b.

When two SPMTs in the same pixel are hit in the event time window, the charges are summed and the lowest of the hit-time is chosen. The time window depends on the datasets and are detailed in Section 4.1.2. The SPMTs being located close to each other, we expect the time difference between two successive physics signals, two photons being collected, to be small. The first hit time is chosen because it can be considered as the relative propagation time of the photons that went the "straightest", i.e. that went under the less perturbation of the two. The timing is thus more representative of the event location.

The only potential problem in using this first time come from the Dark Noise (DN). Its time distribution is uniform over the signal and could come before a physics signal on the other SPMT in the pixel. In that case, the time information in the pixel become irrelevant and we lose the timing information for this part of the detector. As illustrated in Figure 4.2b the image dimension have been optimized so that at most two SPMTs are in the same pixel while keeping the number of empty pixels relatively low to prevent this kind of issue.

While it could be possible to use larger images (more pixel) to prevent overlapping, keeping image small images gives multiple advantages:

- As presented in Section 4.1.1, the convolution filter we use are  $3 \times 3$  convolution filter, meaning that if SPMTs would be separated by more than one pixel, the first filter would only see one SPMT per filter. This behavior would be kind of counterproductive as the first convolution block would basically be a transmission layer and would just induce noise in the data.
- It keep the network relatively small, while this do not impact the convolution layers, the flatten operation just before the FCDNN make the number parameters in the first layer of it dependent on the size of the image.
- It reduce the number of empty pixel in the image.

The question of empty pixel is an important question in this data representation. There is two kind of empty pixels in the data.

The first kind is pixel that contain a SPMT but the SPMT did not get hit nor registered any dark noise during the event. In this case, the charge channel is zero, which have a physical meaning but then come the question of the time layer. One could argue that the correct time would be infinity (or the largest number our memory allows us) because the hit "never" happened, so extremely far from the time of the event. This cause numerical problem as large number, in the linear operation that are happening in the convolution layers, are more significant than smaller value. We could try to encode this feature in another way but no number have any significance due to our time being relative to the trigger of the experiment so  $-1$  for example is out of question. Float and Double gives us access to special value such as NaN (Not a Number) [125] but the behavior is to propagate the NaN which leaves us with NaN for energy and position. We choose to keep the value 0 because it's the absorbing element of multiplication, absorbing the "information" of the parameter it would be

1960 multiplied by. It also can be thought as no activation in the ReLU activation function. It's important  
1961 to keep in mind the fact that a part of the detector that has not been hit is also an information: There  
1962 is no signal in this part of the detector. This problematic will be explored in more details in Chapter  
1963 5.

1964 The second kind of pixels are the ones that do not represent parts of the detector such as the corners  
1965 of the image. The question is basically the same, what to put in the charge and the time channel. The  
1966 decision is to set the charge and time to 0 following the above reasoning.

1967 Another problematic that happens with this representation, and this is not dependent of the chosen  
1968 projection, is the deformation in the edges of the image and the loss of the neighbouring information  
1969 in the for the SPMTs at the edge of the image  $\phi \sim 180^\circ$ . This deformation and neighbouring loss  
1970 could be partially circumvented as explained in Section 4.4

### 1971 4.1.3 Dataset

1972 In this study we will discuss two datasets of one millions prompt signal of IBD events.

#### 1973 J21

1974 The first one comes from the JUNO official MC simulation J21v1r0-Pre2 (released the 18th August  
1975 2021). This historical version is the one on which the classical SPMT reconstruction algorithm was  
1976 developed. This classical methods is based on the time likelihood presented Section 3.3 for the vertex  
1977 reconstruction, and compute the energy by correcting the detector effect on the ratio  $N_{pe}/E_{dep}$ . It is  
1978 detailed in Chapter 4 of [34]. This dataset is used as a reference for comparison to classical algorithm  
1979 performances. The data in this dataset is *detsim* level (see Section 2.6) which includes no digitization,  
1980 no DAQ and therefore no reconstruction of PMT signals. Only the number of PEs that hit a PMT and  
1981 the hit times are provided. A fast simulation based on gaussian drawings produces charges, with  
1982 bias and variability, and the equivalent for times. The drawings parameters were adjusted based on  
1983 [81, 126]. Because there is no charge reconstruction, the timing on the event is based on the Geant4  
1984 simulation, and so  $t = 0$  is the moment the positron is created in the CD. To prevent correlation  
1985 between the numerical value of the time of the first hit  $t_0$  and the radius of the event, we offset all  
1986 time by this first hit time. Without simulation of the charge reconstruction, we cannot simulate the  
1987 event trigger, we thus add an arbitrary time cut at a  $t_0 + 1000$  ns.

#### 1988 J23

1989 The second comes from the JUNO official monte-carlo simulations J23.0.1-rc8.dc1 (released the 7th  
1990 January 2024). The data is *calib* level (see Section 2.6). Here the charge comes from the waveform  
1991 integration, the time window resolution and trigger decision are all simulated inside the software.

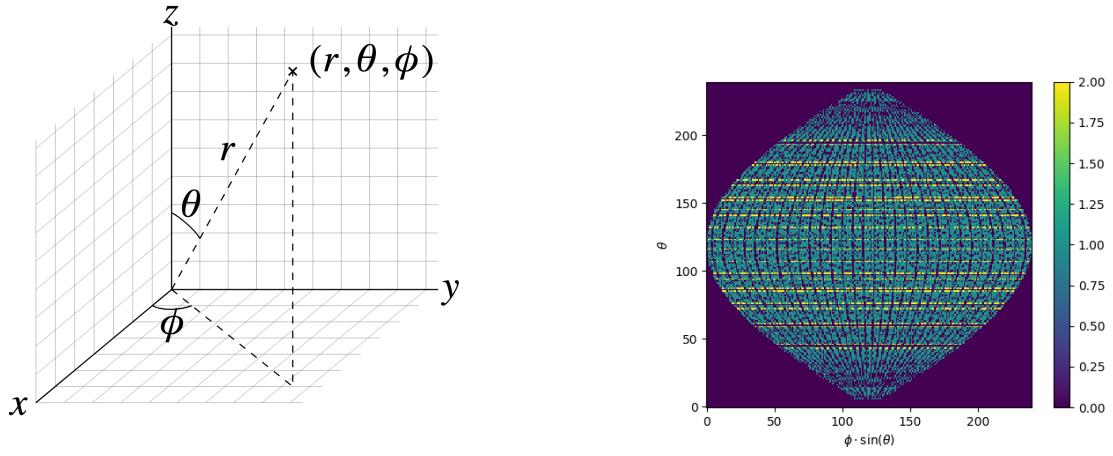
1992 To put in perspective this amount of data, the expected IBD rate in JUNO is 47 / days. Taking into  
1993 account the calibration time, and the source reactor shutdown, it amounts to  $\sim 94'000$  IBD events  
1994 in 6 years. With this million of events, we are training the equivalent of  $\sim 10$  years of data. With  
1995 this amount we reach a density of  $4783 \frac{\text{event}}{\text{m}^3 \cdot \text{MeV}}$ , meaning our dataset is representative of the multiple  
1996 event scenarios that could be happening in the detector.

1997 While we expect and hope the MC simulation to give a realistic representation of the detector,  
1998 there could be effects, even after the fine-tuning on calibration data, that the simulation cannot handle.  
1999 Thus, once the calibration will be available, we will need to evaluate, and if needed retrain, the  
2000 network on calibration data to establish definitive performances.

The simulated data is composed of positron events, uniformly distributed in the CD volume and in kinetic energy over  $E_k \in [0; 9]$  MeV producing a deposited energy  $E_{dep} \in [1.022; 10.022]$  MeV. This is done to mimic the signal produced by the IBD prompt signal. Uniform distributions are used so that the CNN does not learn a potential energy distribution, favoring some part of the energy spectrum instead of other.

#### 4.1.4 Data characteristics

To delve a bit into the kind of data we will use, you can find in Figure 4.2b the repartition of the SPMTs in the image. The color represent the number of SPMTs per pixel.



(A) Spherical coordinate system used in JUNO for reconstruction

(B) Repartition of SPMTs in the image projection. The color scale is the number of SPMTs per pixel

FIGURE 4.2

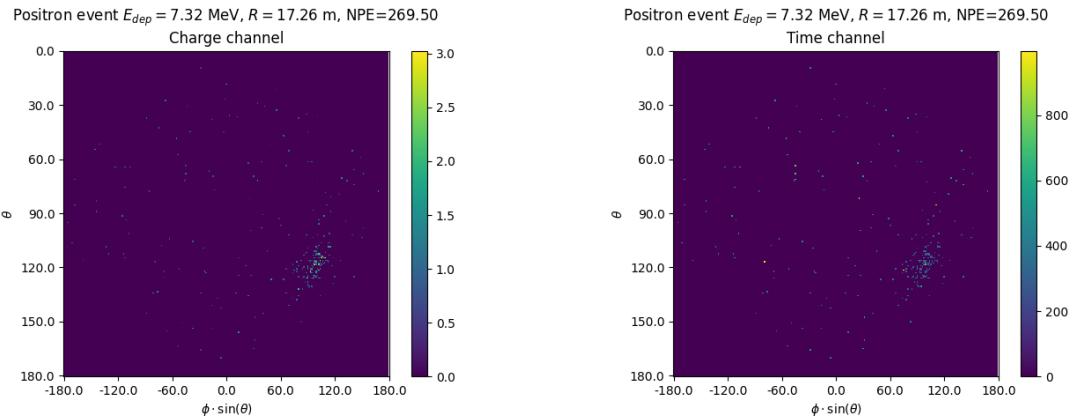


FIGURE 4.3 – Example of a high energy, radial event. We see a concentration of the charge on the bottom right of the image, clear indication of a high radius event. **On the left:** the charge channel. The color is the charge in each pixel in NPE equivalent. **On the right:** The time channel in nanoseconds.

See also Figures 4.3 to 4.6 - and the explanation in their captions - which present events from J23 for different positions and energies. We see some characteristics and we can instinctively understand how the CNN could discriminate different situations.

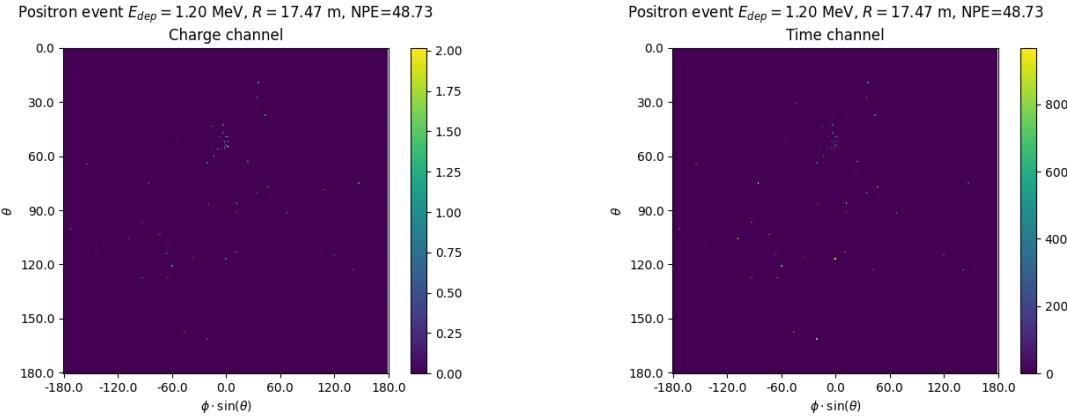


FIGURE 4.4 – Example of a low energy, radial event. The signal here is way less explicit, we can kind of guess that the event is located in the top middle of the image.  
**On the left:** the charge channel. The color is the charge in each pixel in NPE equivalent.  
**On the right:** The time channel in nanoseconds.

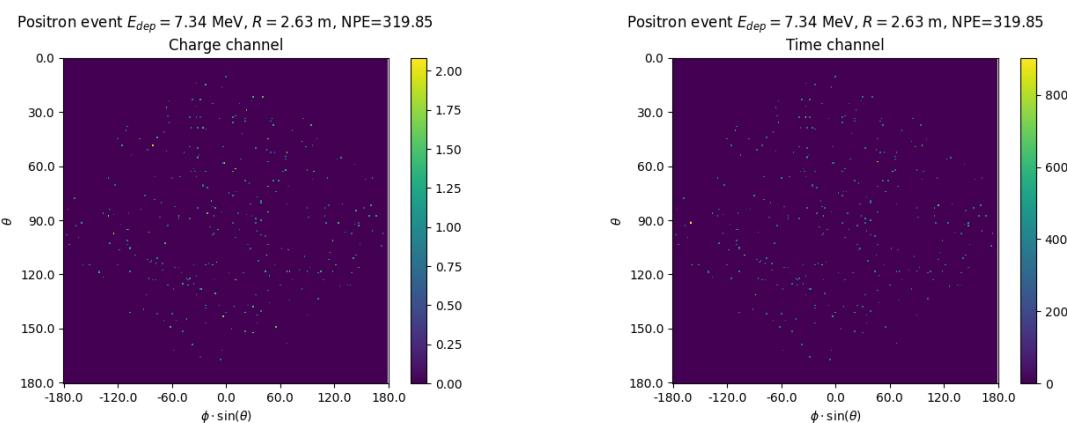


FIGURE 4.5 – Example of a high energy, central event. In this image we can see a lot of signal but uniformly spread, this is indicative of a central event.  
**On the left:** the charge channel. The color is the charge in each pixel in NPE equivalent.  
**On the right:** The time channel in nanoseconds.

To give an idea of the strength of the signal in comparison to the dark noise background, Figure 4.7a present the distribution of the ratio of NPE per deposited energy. Assuming a linear response of the LS we can model:

$$NPE_{tot} = E_{dep} \cdot P_{mev} + D_N \quad (4.5)$$

$$\frac{NPE_{tot}}{E_{dep}} = P_{mev} + \frac{D_N}{E_{dep}} \quad (4.6)$$

where  $NPE_{tot}$  is the total number of PE detected by the event,  $P_{mev}$  is the mean number of PE detected per MeV and  $D_N$  is the dark noise contribution that is considered energy independent. In the case where the readout time window is dependent of the energy the dark noise contribution become energy dependant, also the LS response is realistically energy dependant but Figure 4.7a shows that we are heavily dominated by the stochastic behavior of light emission and detection.

The fit shows a light yield of 40.78 PE/MeV and a dark noise contribution of 4.29 NPE. As shown in

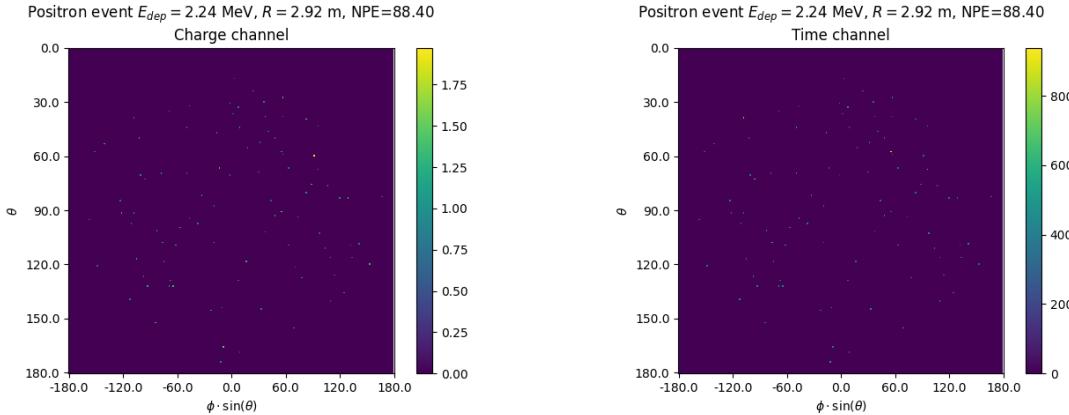


FIGURE 4.6 – Example of a low energy, central event. Here there is no clear signal, the uniformity of the distribution should make it central. **On the left:** the charge channel. The color is the charge in each pixel in NPE equivalent. **On the right:** The time channel in nanoseconds.

2018    Figure 4.7b, the physics makes for 90% of the signal at low energy.

## 2019    4.2 Training

2020    The optimizer used for the training is the Adam [95] optimizer, with a learning rate  $\lambda$  of 1e-3. The  
2021    other hyperparameters were left to their default value ( $\beta_1 = 0.9$ ,  $\beta_2 = 0.999$  and  $\epsilon = 1e^{-8}$ ). The  
2022    learning rate was reduced exponentially during the training at a rate of  $\gamma = 0.95$ , thus  $\lambda_{i+1} = 0.95\lambda_i$   
2023    where  $i$  is the epoch.

2024    Following the lifecycle presented in Section 3.1.3, the training used a batch size of 64 events meaning  
2025    that, each step, the loss is computed on 64 events before updating the NN parameters. An epoch is  
2026    composed of 10k steps, thus each epoch, the NN sees 640k events. The training last for 30 epochs, so  
2027    overall the NN goes through 19.2 millions events or 19.2 times the dataset.

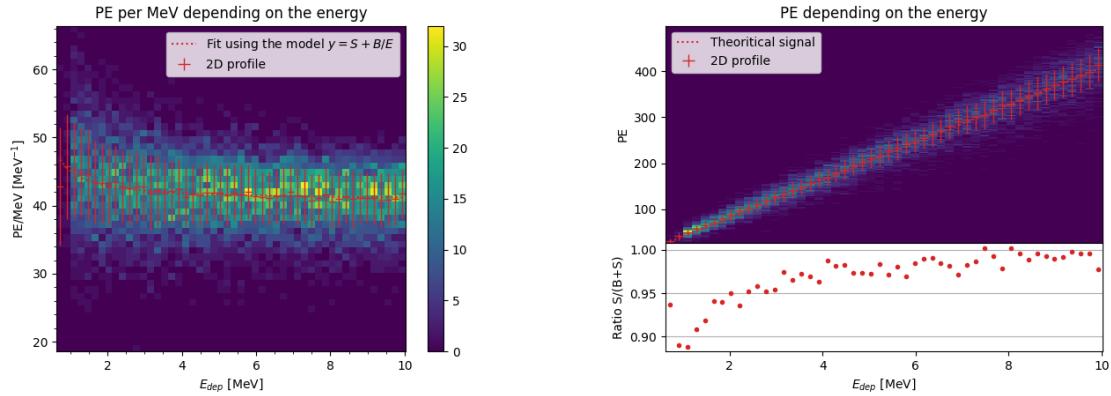
2028    The number of epoch, batch size, learning rate and its decay were fine-tuned during the development  
2029    of the CNN.

## 2030    4.3 Results

2031    Before presenting the results, let's discuss the different observables.

2032    The events are considered point-like in this study. The target truth position, or vertex, is the mean  
2033    position of the energy deposits of the positron and the two annihilation gammas. This approximation  
2034    for point-like interaction is also used for the likelihood study presented in Section 3.3 and in previous  
2035    ML studies presented in section 3.3.3 [114].

2036    Due to the symmetries of the detector, we mainly consider and discuss the bias and precision evolution  
2037    depending on the radius  $R$  but we will still monitor the performances depending on the spherical  
2038    angle  $\theta$  and  $\phi$ . From the detector construction and effect we expect dependency in radius due to the  
2039    TR area effect presented in Section 3.3 and the possibility for the positron or the gammas to escape  
2040    from the CD for positrons interacting near the edge. We also expect dependency on  $\theta$ , the top of the  
2041    experiment being non-instrumented due to the filling chimney. It is also to be noted that the events



(A) Distribution of PE/MeV in the J23 Dataset. This distribution is profiled and fitted using equation 4.6

(B) On top: Distribution of PE vs Energy. On bottom: Using the values extracted in 4.7a, we calculate the ration signal over background + signal

FIGURE 4.7

in the dataset are uniformly distributed in the CD, and so are uniformly distributed in  $R^3$  and  $\phi$ . The  $\theta$  distribution is not uniform and we will have more event for  $\theta \sim 90^\circ$  than  $\theta \sim 0^\circ$  or  $\theta \sim 180^\circ$ .

We define multiple energy in JUNO:

- $E_\nu$ : The energy of the neutrino.
- $E_k$ : The kinetic energy of the resulting positron from the IBD.
- $E_{dep}$ : The deposited energy of the positron and the two annihilation gammas.
- $E_{vis}$ : The equivalent visible energy, so  $E_{dep}$  after the detector effect such as the LS response non-linearity.
- $E_{rec}$ : The reconstructed energy by the reconstruction algorithm. The expected value depend on the algorithm we discuss about. For example the algorithm presented in Section 3.3 reconstruct  $E_{vis}$  while the ones presented in section 3.3.3 reconstruct  $E_{dep}$ .

In this study, we will set  $E_{dep}$  as our target for energy reconstruction. This choice is motivated by the ease with which we can retrieve this information in the monte-carlo data while  $E_{vis}$  is less trivial to retrieve.

### 4.3.1 J21 results

The best results comes from the Gen<sub>30</sub> model, meaning then 30th model generated using the table 4.1: Gen<sub>30</sub>:  $N_{blocks} = 3$ ,  $N_{channels} = 32$ , FCDNN configuration: 2048 \* 2 + 1024 \* 2, Loss  $\equiv E + V$ .

The performances of its reconstruction are presented in blue in Figure 4.8. Superimposed in black is the performances of the classical algorithm from [34].

#### Energy reconstruction

By looking at the Figure 4.8a and 4.8b, the CNN has similar performances in its energy resolution. Important biases, however, appear at low and high energy.

This is explained by looking at the true and reconstructed energy distributions in Figure 4.10a. We see that the distributions are similar for energies before 8 MeV but there is an excess of event

reconstructed with energies around 9 MeV while a lack of them for 10 MeV. The neural network seems to learn the energy distribution and learn that it exist almost no event with an energy inferior to 1.022 MeV and not event with an energy superior to 10 MeV.

The first observation is a physics phenomena: for a positron, its minimum deposited energy is the mass energy coming from its annihilation with an electron 1.022 MeV. There is a few event with energies inferior to 1.022 MeV, in those case the annihilation gammas or even the positron escape the detector. The deposited energy in the LS is thus only a fraction of the energy of the event.

The second observation is indeed true in this dataset but has no physical meaning, it is an arbitrary limit because the physics region of interest is mainly between 1 and 9 MeV of deposited energy (Figure 2.2). By learning the energy distribution, the CNN pull event from the border of it to more central value. That's why the energy resolution is better: the events are pulled in a small energy region , thus a small variance but the bias become very high (Figure 4.8a).

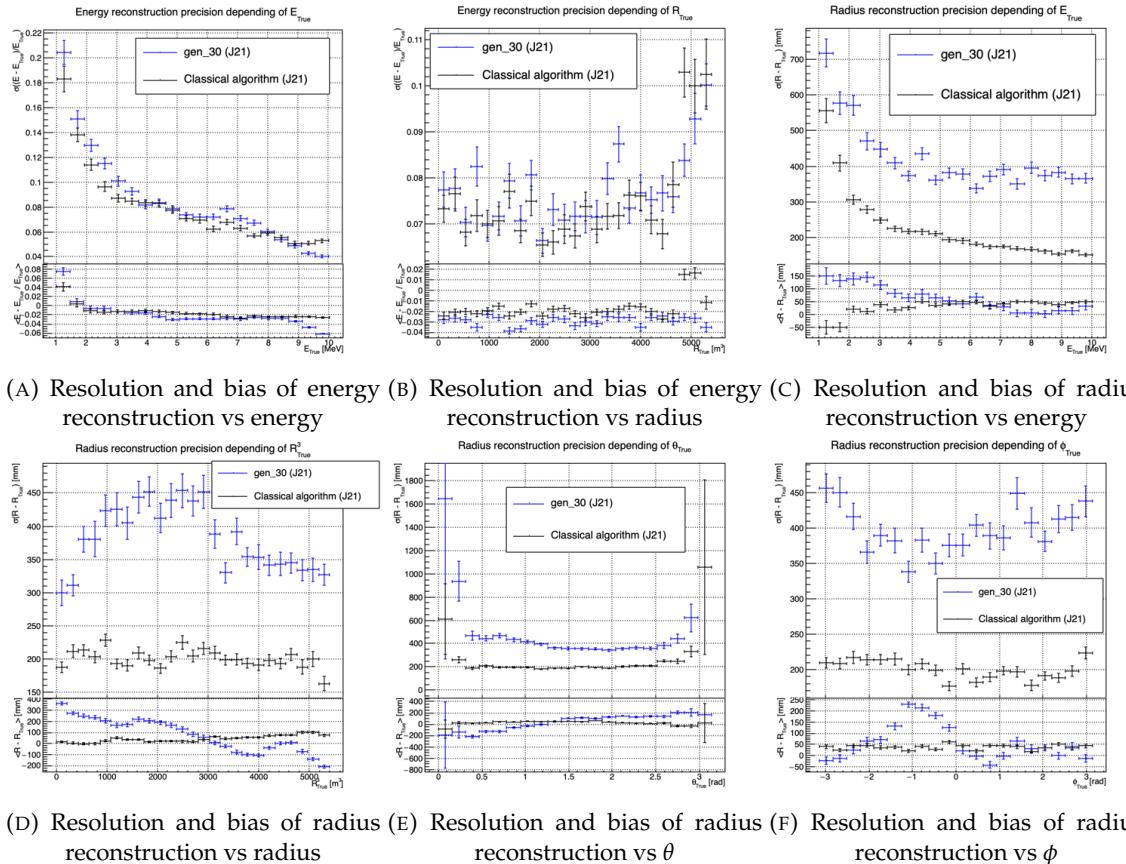


FIGURE 4.8 – Reconstruction performance of the Gen<sub>30</sub> model on J21 data and its comparison to the performances of the classic algorithm “Classical algorithm” from [34]. The top part of each plot is the resolution and the bottom part is the bias.

This behavior also explain the heavy bias at low energy in Figure 4.8a. The energy bias of the CNN is fairly constant over the energy range, it is interesting to note that the energy bias depending on the radius is a bit worse than the classical method.

2081 **Vertex reconstruction**

2082 For the vertex reconstruction we do not study  $x$ ,  $y$  and  $z$  independently but we use  $R$  as a proxy  
 2083 observable. Figure 4.9 shows the residual distribution of the different vertex coordinates. We see  
 2084 that  $R$  errors and biases are slightly superior to the cartesian coordinates, thus  $R$  is a conservative  
 2085 proxy observable to discuss the subject of vertex reconstruction.

2086 The comparison of radius reconstruction between the classical algorithm and Gen<sub>30</sub> are presented in  
 2087 the Figures 4.8c, 4.8d, 4.8e and 4.8f. The resolution obtained by the CNN is twice worse in average,  
 2088 and worse in all studied regions. In energy, Figure 4.8c, where we see a degradation of almost 20cm  
 2089 over the energy range. When looking over the true event radius, Figure 4.8d, we lose between 30  
 2090 and 45cm of resolution. The performances are the best for central and radial event.

2091 The precision also worsen when looking at the edge of the image  $\theta \approx 0, \theta \approx 2\pi$  respectively the top  
 2092 and bottom of the image, and when  $\phi \approx -\pi$  and  $\phi \approx \pi$  respectively the left and right side of the  
 2093 image.

2094 The bias in radius reconstruction is about the same order of magnitude depending of the energy but  
 2095 is of opposite sign. As for the energy, this behavior is studied in more details in Section 4.3.2. Over  
 2096 radius,  $\theta$  and  $\phi$  the bias is inconsistent, sometimes event better than the classical reconstruction but  
 2097 can also be much worse than the classical method. This could come from the specialisation of some  
 2098 filters in the convolutional layers for specific part of the detector that would still work “correctly” for  
 2099 other parts but with much less precision.

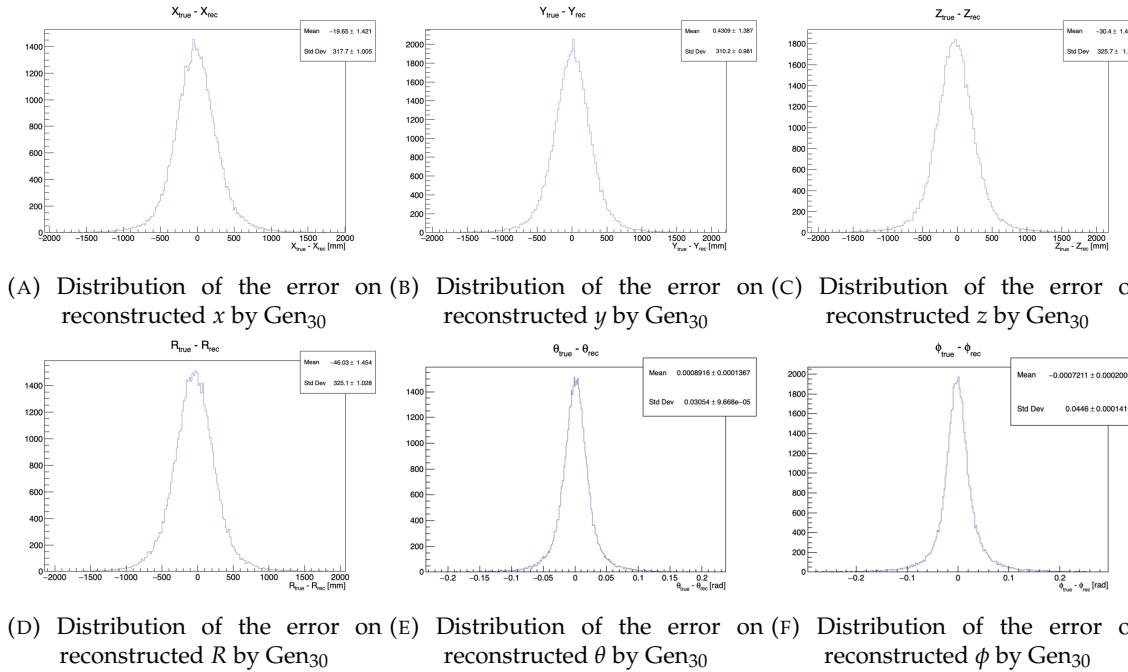
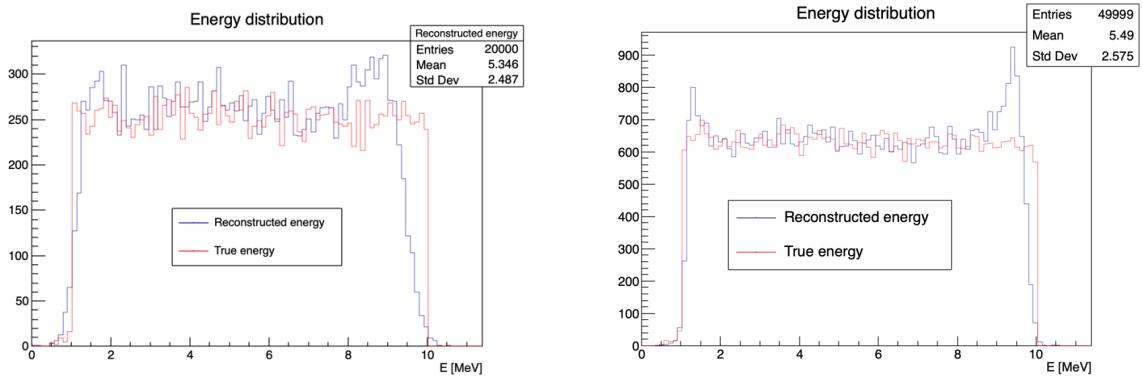


FIGURE 4.9 – Residual distribution of the different component of the vertex by Gen<sub>30</sub>.  
 The reconstructed component are  $x$ ,  $y$  and  $z$  but we see similar behavior in the error of  
 $R$ ,  $\theta$  and  $\phi$ .

2100 As mentioned in the introduction of this chapter, this CNN initially served as a tool for learning  
 2101 about machine learning and JUNO’s detector and software. It eventually became necessary for use  
 2102 as an SPMT reconstruction tool in Chapter 7, so we made some optimizations. However, we did not  
 2103 invest much time in fully addressing its issues.



(A) Distribution of Gen<sub>30</sub> reconstructed energy and true energy of the analysis dataset (J21)      (B) Distribution of Gen<sub>42</sub> reconstructed energy and true energy of the analysis dataset (J23)

FIGURE 4.10

### 2104 4.3.2 J21 Combination of classic and ML estimator

As it has been presented in previous section, there are instances where the reconstructed energy and vertex behaves differently between the neural network and the classic algorithm. For instance, if we look at Figure 4.8c, we see that while the CNN tend to overestimate the radius at low energy while the classical algorithm seems to underestimate it. Let's designate the two reconstruction algorithms as estimator of  $X$ , the truth about the event in the phase space  $(E, x, y, z)$ . The CNN and the classical algorithm are respectively designated as  $\theta_N(X)$  and  $\theta_C(X)$ .

$$E[\theta_N] = \mu_N + X; \text{Var}[\theta_N] = \sigma_N^2 \quad (4.7)$$

$$E[\theta_C] = \mu_C + X; \text{Var}[\theta_C] = \sigma_C^2 \quad (4.8)$$

2105 where  $\mu$  is the bias of the estimator and  $\sigma^2$  its variance.

2106 Now if we were to combine the two estimators using a simple mean

$$\hat{\theta}(X) = \frac{1}{2}(\theta_N(X) + \theta_C(X)) \quad (4.9)$$

then the variance and mean would follow

$$E[\hat{\theta}] = \frac{1}{2}E[\theta_N] + \frac{1}{2}E[\theta_C] \quad (4.10)$$

$$= \frac{1}{2}(\mu_N + X + \mu_C + X) \quad (4.11)$$

$$= \frac{1}{2}(\mu_N + \mu_C) + X \quad (4.12)$$

$$\text{Var}[\hat{\theta}] = \frac{1}{4}\sigma_N^2 + \frac{1}{4}\sigma_C^2 + 2 \cdot \frac{1}{4} \cdot \sigma_{NC} \quad (4.13)$$

$$= \frac{1}{4}\sigma_N^2 + \frac{1}{4}\sigma_C^2 + \frac{1}{2} \cdot \sigma_{NC} \quad (4.14)$$

$$= \frac{1}{4}\sigma_N^2 + \frac{1}{4}\sigma_C^2 + \frac{1}{2} \cdot \sigma_N \sigma_C \rho_{NC} \quad (4.15)$$

2107 Where  $\sigma_{NC}$  is the covariance between  $\theta_N$  and  $\theta_C$  and  $\rho_{NC}$  their correlation.

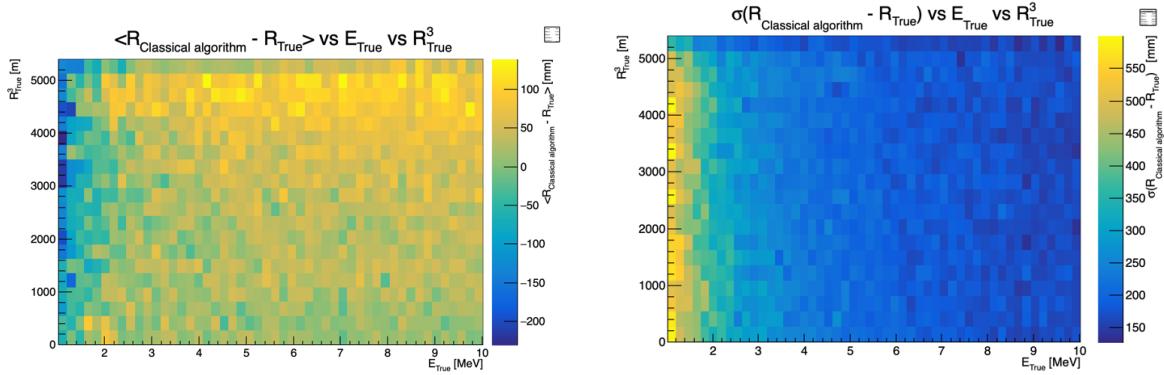


FIGURE 4.11 – Radius bias (on the left) and resolution (on the right) of the classical algorithm in a  $E, R^3$  grid

2108 We see immediately that if the two estimators are of opposite bias, the bias of the resulting estimator  
 2109 is reduced. For the variance, it depends of  $\rho_{NC}$  but in this case if  $\sigma_C^2$  is close to  $\sigma_N^2$  then even for  
 2110  $\rho_{NC} \lesssim 1$  then we can gain in resolution.

2111 By generalising the equation 4.9 to

$$\hat{\theta}(X) = \alpha\theta_N + (1 - \alpha)\theta_C; \alpha \in [0, 1] \quad (4.16)$$

2112 we can determine an optimal  $\alpha$  for two combined estimators. The estimators with the smallest  
 2113 variance

$$\alpha = \frac{\sigma_C^2 - \sigma_N\sigma_C\rho_{NC}}{\sigma_N^2 + \sigma_C^2 - 2\sigma_N\sigma_C\rho_{NC}} \quad (4.17)$$

2114 and the estimator without bias

$$\alpha = \frac{\mu_C}{\mu_C - \mu_N} \quad (4.18)$$

2115 See annex A for demonstration.

2116 We present in this section the result of the estimator with the smallest variance.

2117 Its pretty clear from the results shown in Figure 4.8 that the bias, variances and correlation are not  
 2118 constant across the  $(E, R^3)$  phase space. We thus compute those parameters in a grid in  $E$  and  $R^3$  for  
 2119 the following results as illustrated in 4.11.

2120 The map we are using are composed of 20 bins for  $R^3$  going from 0 to 5400 m<sup>3</sup> (17.54 m) and 50 bins  
 2121 in energy ranging from 1.022 to 10.022 MeV. In the case where we are outside the grid, we use the  
 2122 closest cell.

2123 The performance of this weighted mean is presented in Figure 4.12. We can see that even when the  
 2124 CNN resolution is much worse than the classical algorithm, it can still bring some information thus  
 2125 improving the resolution. This comes from the correlation of the reconstruction error to be smaller  
 2126 than 1 as presented in Figure 4.13. We even see some anticorrelation in the radius reconstruction for  
 2127 High radius, high energy, event.

2128 This technique is not suited for realistic reconstruction, we rely too much on the knowledge of the  
 2129 resolution, bias and correlation between the two methods. While this is possible to determine using  
 2130 simulated data or calibration sources, the real data might differ from our model and we would need  
 2131 to really well understand the behavior of the two system. But this is a good tool to detect that  
 2132 algorithms don't all use the same information, and is a first step to identify new information that  
 2133 could be brought to the best algorithms, to improve their performance.

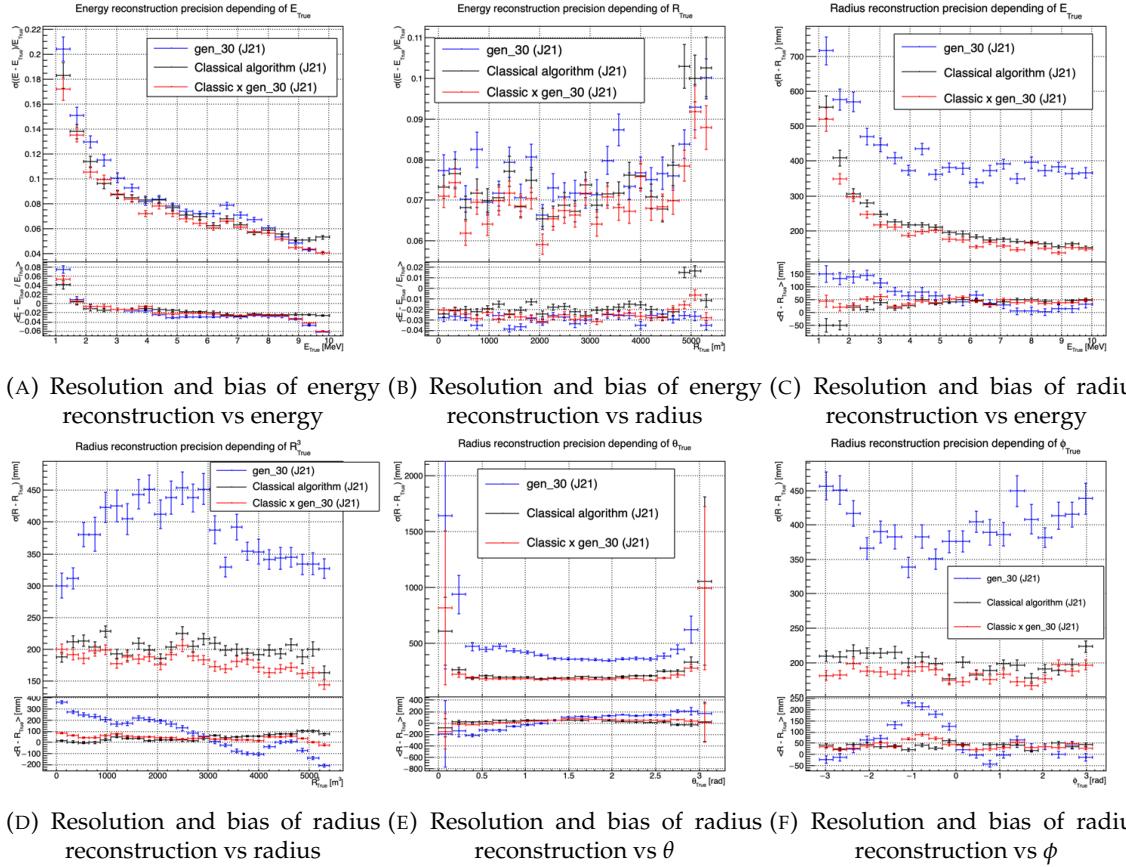


FIGURE 4.12 – Reconstruction performance of the Gen30 model on J21, the classic algorithm “Classical algorithm” from [34] and the combination of both using weighted mean. The top part of each plot is the resolution and the bottom part is the bias.

### 4.3.3 J23 results

We needed for Chapter 7 a SPMT reconstruction tool to run the comparison with LPMT. We thus retrained the SPMT CNN on newer, more realistic data.

The J21 simulation is fairly old and newer version, such as J23, include refined measurements of the light yield, reflection indices of materials of the detector, structural elements such as the connecting structure and more realistic dark noise. Additionally, the trigger, waveform integration and time window are defined using the algorithms that will ultimately be used by the collaboration to process real physics events.

We retrained the models defined in 4.1.1 on the J23 data and used the same hyperparameter optimisation procedure. The results from the best architecture, Gen<sub>42</sub>, are presented in Figure 4.14. Following the table 4.1, Gen<sub>42</sub>:  $N_{blocks} = 3$ ,  $N_{channels} = 64$ , FCDNN configuration:  $4096 * 2$ , Loss  $\equiv E + V$ .

#### Energy reconstruction

The results of the energy reconstruction are presented in Figures 4.14a and 4.14b. The resolution is close to the one of the classical algorithm with the exception of the start and end of the spectrum.

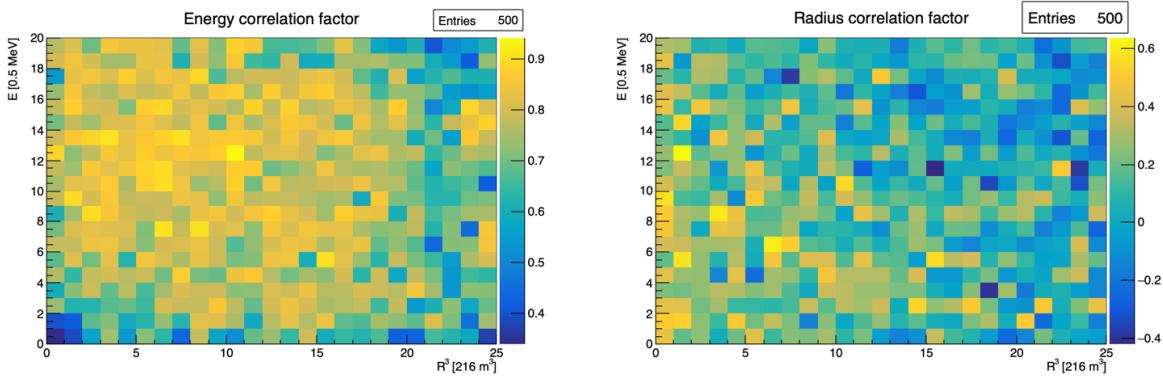


FIGURE 4.13 – Correlation between CNN and classical method reconstruction (on the left) for energy and (on the right) for radius in a  $E, R^3$  grid

2149 This is the same effect that we saw with Gen<sub>30</sub>, events are pulled from the edge of the distribution,  
 2150 resulting in smaller resolution but heavy biases.

### 2151 Vertex reconstruction

2152 The vertex reconstruction, presented in Figures 4.14c, 4.14d, 4.14e and 4.14f is not yet to the level of  
 2153 the classical reconstruction but the degradation is smaller than for Gen<sub>30</sub> being at most a difference  
 2154 of 15cm of resolution and closing to the performance of the classical algorithm in the most favourable  
 2155 condition. Gen<sub>42</sub> has also very little bias in comparison with the classical method with the exception  
 2156 of the transition to the TR area and at the very edge of the detector.

2157 With a more realistic description of the propagation and collection of scintillation photons, of the charge  
 2158 and time resolutions, of the DN and of the trigger, it seems new features can be identified by  
 2159 the CNN.

2160 Unfortunately could not rerun the classical algorithm over the J23 data, as the algorithm was optimised  
 2161 for J21 and was not included and maintained over J23. The combination method need for  
 2162 the two estimators to be run on the same set of event, which was impossible without the classical  
 2163 algorithm being maintained for J23.

## 2164 4.4 Conclusion and prospect

2165 In this chapter we have developed a CNN for the reconstruction of IBD prompt signals. This work  
 2166 was the opportunity to learn about machine learning and neural networks, and familiarise ourselves  
 2167 with JUNO's detector and software.

2168 This work was revisited for the needs of Chapter 7, providing a reconstruction tools for the SPMT.

2169 The CNN we developed suffers limitations in its performance. We think one of the reasons for this  
 2170 lies in the data representation. First, a lot of training time and resources is consumed going and  
 2171 optimizing over pixels with no physical meaning, notably the time information in case of no hit.  
 2172 This problem origin from the planar projection and is also a specificity of the SPMT system, where  
 2173 a low number of PMT fire per event resulting in empty pixels. To overcome this problematic, i.e.  
 2174 what is the time of a PMT that was never hit, we could transform this channel into a dimension. This  
 2175 would results in an image with multiple charge channels, each one representing the charge sum in a  
 2176 time interval.

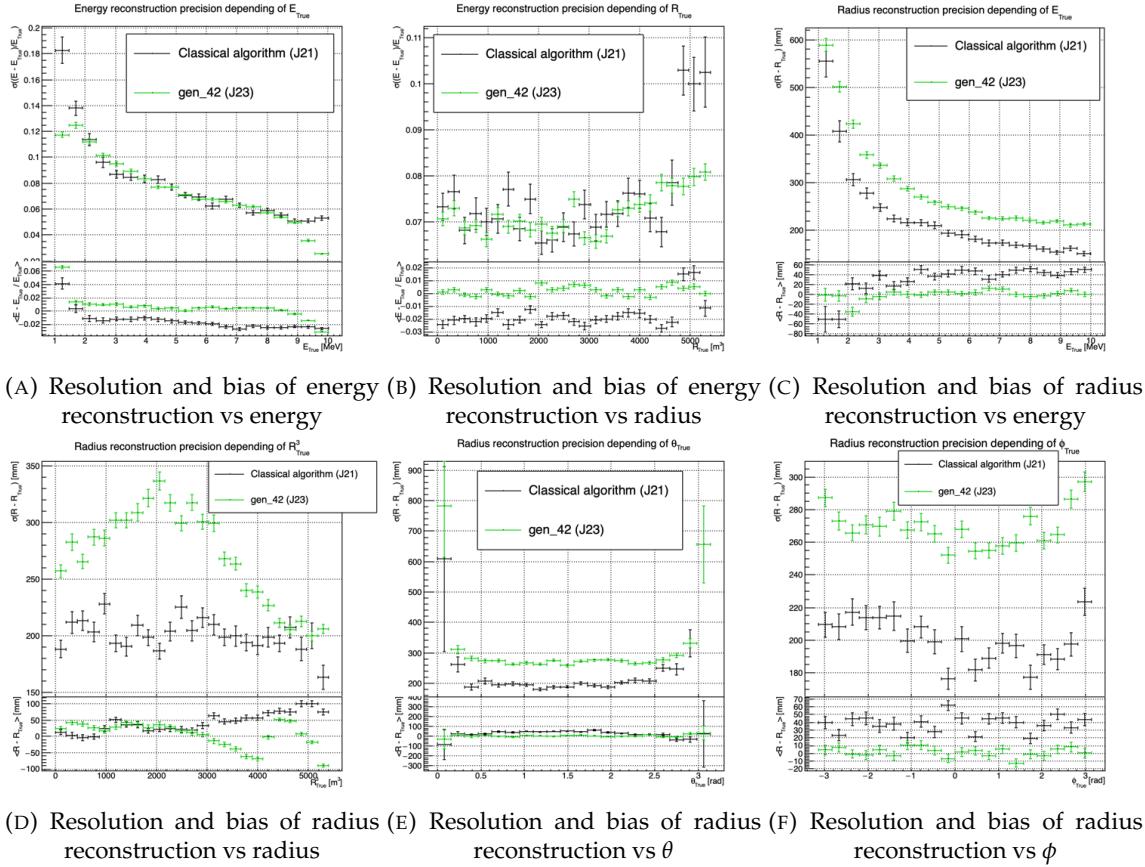


FIGURE 4.14 – Reconstruction performance of the Gen42 model on J23 data and its comparison to the performances of the classic algorithm “Classical algorithm” from [34]. The top part of each plot is the resolution and the bottom part is the bias.

Even the best CNN design should at some point hit another limitation : the necessity to project the spherical image on a sphere. It would then need to optimize itself to take into account edges cases such as event at the edge of the image and deformation of the charge distribution. we could imagine a two part CNN where the first part reconstruct the  $\theta$  and  $\phi$  spherical coordinates and then rotate the image to locate the event in the center of the image. The second part, from this rotated image, would reconstruct the radius and energy of the event. Another possibility is to use a kind of algorithm that does not impose a planar projection, like a GNN. It has other advantages, as will be presented in the next chapter, where we propose a GNN to reconstruct IBD’s with the LPMT and SPMT systems.

The CNN we developed suffers limitations in its performance. We think one of the reasons for this lies in the data representation. A lot of training time and resources is consumed going and optimizing over pixel with no physical meaning, the NN needs to optimized itself to take into account edges cases such as event at the edge of the image and deformation of the charge distribution.

Those problems could be circumvented, we could imagine a two part CNN where the first part reconstruct the  $\theta$  and  $\phi$  spherical coordinates and then rotate the image to locate the event in the center of the image. The second part, from this rotated image, would reconstruct the radius and energy of the event.

To overcome the time problematic, i.e. what is the time of a PMT that was never hit, we could transform this channel into a dimension. This would results in an image with multiple charge channels, each one representing the charge sum in a time interval.

<sup>2196</sup> Another possibility is to use a kind of algorithm that does not impose a planar projection, like a  
<sup>2197</sup> GNN. It has other advantages, as will be presented in the next chapter, where we propose a GNN to  
<sup>2198</sup> reconstruct IBD's with the LPMT system.

<sup>2199</sup> **Chapter 5**

<sup>2200</sup> **Graph representation of JUNO for  
IBD reconstruction**

<sup>2202</sup> "The Answer to the Great Question of Life, the Universe and  
Everything is Forty-two"

*Douglas Adams, The Hitchhiker's Guide to the Galaxy*

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<sup>2217</sup> In Section 3.3.3, we showed that all ML methods developed before this thesis to reconstruct IBDs have similar results, and that their performance is very similar to that of the classical, likelihood-based algorithm. We think these similarities can reasonably be explained by this: the input data used by all these methods to compute  $E$  or  $\vec{X}$  is the same full list of PMT integrated signals  $\{(Q_i, t_i); i \in 1, \dots, N_{PMTs}\}$ , and by the high level of sophistication of the detector's description in the likelihood. It's probable that the likelihood method looses very little information.

<sup>2224</sup> May be some was, but that the ML algorithms were not designed well enough to recover it. It's also reasonable to think that ML algorithms will make a difference when, instead of the list of  $(Q_i, t_i)$ , a rawer information will be used in input, like the full waveform. To actually be able to learn from such a complex and high dimensional input, well designed architectures (that would guide the learning toward the solution) are necessary. In any case, it seemed welcome to us to propose an additional algorithm, with an original architecture.

<sup>2230</sup> For the fist stage of its development, the purpose of this part of my thesis, we considered it was enough to also take the  $(Q_i, t_i)$  list as the input. While achieving equivalent performance with simpler input might suggest that the architecture is not immediately advantageous, it remains crucial to explore the performance with more complex, rawer inputs such as full waveforms. This is where the true potential of the architecture could emerge, as it could better capture the intricacies that simpler inputs fail to represent. If performance does not improve with these richer inputs, it would then be appropriate to question the relevance of this approach.

The algorithm we propose is a GNN. It also has the advantage of addressing sphericity issues described in Chapter 4. From this graph representation, we can construct a neural network that will process the data while keeping some interesting properties. For example the rotational invariance, i.e. the energy and radius of the event do change by rotation our referential. For more details see Section 3.2.3. Graph representation also has the advantage to be able to encode global and higher order informations.

## 5.1 Data representation

In Section 3.3.3, we mentioned a GNN developed before the beginning of this thesis to reconstruct IBD energies in JUNO [114]. In their approach: nodes of the graph correspond to 3072 pixels representing geometric regions of the detector and the information of the  $\sim 6$  LPMTs found in a pixel are then aggregated on those nodes. This aggregation serves to simplify the data input, though at the potential cost of losing finer-grained details. The network then process the data using the equivalent of convolution but on graph [104]. In the first layer, each node is connected only with its direct neighbours.

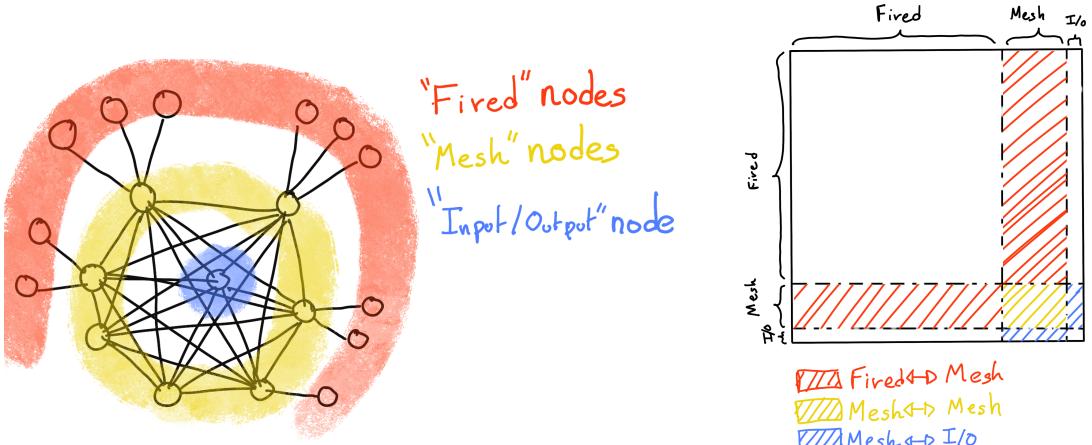
To determine the energy released by an IBD in the LS, it is helpful to determine the position of the main energy deposit. Therefore, relative Q and t's of PMTs all around the sphere is a useful information. If in the first layer only neighbour nodes are linked, several layers are necessary to access this detector-wide information. In an ideal world, we would develop a Graph NN where each PMT is a node (even if it has not been hit in the event under consideration, since this is in itself an information) and where each node is connected to all the other ones. This makes the detector-wide information available as early as the first layer. This architecture might help the network to better learn. Such an architecture can also be motivated this way: one of the strength of GNN's is their capacity to encompass the characteristics of a detector. A node can be the representation of a detector element, and the edge can represent its relationship with other elements. In the case of JUNO, any measurement is collective : an interaction is seen by all the PMTs, with no a priori hierarchy in the role of each. A fully connected GNN is particularly advantageous in JUNO's case, as the lack of a priori hierarchy among the PMTs makes it important to ensure that information is shared globally from the outset. This architecture allows the network to access detector-wide information as early as the first layer, potentially improving learning efficiency. However, this comes at a significant computational cost, which necessitates careful balancing between memory usage and model performance

Another advantage of a GNN is also that it is well adapted to inhomogenous detectors. We therefore tried to build GNNs including both LPMTs and SPMTs.

With 17612 LPMTs and 25600 SPMTs, the ideal fully connected Graph mentioned above is impossible: even excluding self relation and considering the relation to be undirected (the edge from a node A to a node B being the same from as the one from B to A) the amount of necessary edges would be  $n(n - 1)/2$  with  $n = 43212$  nodes. This amounts to 933'616'866 edges. If we encode an information with double precision (64 bits) in what we call an adjacency matrix, illustrated in Figure 3.12, each information we want to encode in the relation would consume 4 GB of data. When adding the overhead due to gradient computation during training, this would put us over the memory capacity of a single V100 gpu card (20 GB of memory). We could use parallel training to distribute the training over multiple GPU but we considered that the technical challenge to deploy this solution was too high.

We finally decided of a middle ground where we define three *families* of nodes:

- The core of the graph is composed of nodes representing geometric regions of the detector. We call those nodes **mesh** nodes. Those mesh nodes are all connected to each other. We keep their number low to gain in memory consumption.



(A) Illustration of the different nodes in our graphs and their relations.

(B) Illustration of what a dense adjacency matrix would look like and the part we are really interested in. Because Fired → Mesh and Mesh → I/O relations are undirected, we only consider in practice the top right part of the matrix for those relations.

FIGURE 5.1

- PMTs in which Photo-Electrons (PE) are found are represented by **fired** nodes. Fired nodes are connected to the mesh node they geometrically belong to.
- A final node is called the input/output node (**I/O**). It is connected to every mesh node. Its features are combinations of signals found in the whole detector.

Those nodes and their relations are illustrated in Figure 5.1a. From this representation, we end up with three distinct adjacency matrix

- A  $N_{\text{fired}} \times N_{\text{mesh}}$  adjacency matrix, representing the relations between fired and mesh. Those relations are undirected.
- A  $N_{\text{mesh}} \times N_{\text{mesh}}$  adjacency matrix, representing the relation between meshes. Those relation are directed.
- A  $N_{\text{mesh}} \times 1$  adjacency between the mesh and I/O nodes. Those relations are undirected.

The adjacency matrix representing those relation is illustrated in Figure 5.1b.

The mesh segmentation is following the Healpix segmentation [127]. This segmentation offer the advantage that almost each mesh have the same number of direct neighbours and it guarantee that each mesh represent the same extent of the detector surface. The segmentation can be infinitely subdivided to provide smaller and smaller pixels. The number of pixel follow the order  $n$  with  $N_{\text{pix}} = 12 \cdot 4^n$ . This segmentation is illustrated in Figure 5.2. To keep the number of mesh small, we use the segmentation of order 2,  $N_{\text{pix}} = 12 \cdot 4^2 = 192$ .

We decided on having the different kind of nodes **mesh (M)**, **fired (F)** and **I/O** have different set of features. The features used in the graph are presented in tables 5.1 and 5.2. Most of the features are low level informations such as the charge or time information but we include some high order features such as

1.  $P_l^h$ : Is the normalized power of the  $l$ th spherical harmonic. For more details about spherical harmonics in JUNO, see annex B.



FIGURE 5.2 – Illustration of the Healpix segmentation. **On the left:** A segmentation of order 0. **On the right:** A segmentation of order 1

2. **A** and **B** are informations that are related the likeliness of the interaction vertex to be on the segment between the center of two meshes.

$$\mathbb{A}_{ij} = (\vec{j} - \vec{i}) \cdot \frac{l_1}{D_{ij}} + \vec{i} \quad (5.1)$$

$$\mathbb{B}_{ij} = \frac{Q_i}{Q_j} \left( \frac{l_2}{l_1} \right)^2 \quad (5.2)$$

$$l_1 = \frac{1}{2} (D_{ij} - \Delta t \frac{c}{n}) \quad (5.3)$$

$$l_2 = \frac{1}{2} (D_{ij} + \Delta t \frac{c}{n}) \quad (5.4)$$

where  $\vec{i}$  is the position vector of the mesh  $i$ ,  $D_{ij}$  is the distance between the center of the meshes  $i$  and  $j$ ,  $Q_i$  the sum of charges on the mesh  $i$ ,  $\Delta t = t_i - t_j$  where  $t_i$  the earliest time on the mesh  $i$  and  $n$  the optical index of the LS. **A** is the vertex between center of meshes distance ratio between  $i$  and  $j$  based on the time information. For **B**, the charge ratio evolve with the square of the distance, so the mesh couple with the smallest **B** should be the one with the interaction vertex between its two center.

Fired	Mesh	I/O
$Q$	$\langle Q_m \rangle$	$\langle X \rangle$
$t$	$\sigma Q_m$	$\langle Y \rangle$
$x$	$\min(t_m)$	$\langle Z \rangle$
$y$	$\max(t_m)$	$\sum Q$
LPMT/SPMT: 1/-1	$\sigma t_m$ $X_m$ $Y_m$ $Z_m$	$P_l^h; l \in [0, 8]$

TABLE 5.1 – Features on the nodes of the graph. All charge are in [nPE], time in [ns] and position in [m].

$Q$  and  $t$  are the reconstructed charge and time of the hit PMTs.  $(x, y, z)$  is the position of the PMTs and the last parameter represent the type of the PMT. It's 1 for LPMT and -1 for SPMT

$Q_m$  and  $t_m$  is the set of charges and time of the PMT belonging the mesh  $m$ .  $(X_m, Y_m, Z_m)$  i the position of the center of the geometric region represented by the mesh  $m$

$(\langle X \rangle, \langle Y \rangle, \langle Z \rangle)$  is the position of the charge barycenter,  $\sum Q$  the sum of the collected charge in the detector and  $P_l^h$  is the relative power of the  $l$ th harmonic. See annex B for details.

Fired → Mesh	Mesh ( $m1$ ) → Mesh ( $m2$ )	Mesh → I/O
$x - X_m$	$X_{m1} - X_{m2}$	$\langle X \rangle - X_m$
$y - Y_m$	$Y_{m1} - Y_{m2}$	$\langle Y \rangle - Y_m$
$z - Z_m$	$Z_{m1} - Z_{m2}$	$\langle Z \rangle - Z_m$
$t - \min(t_m)$	$\min(t_{m1}) - \min(t_{m2})$	$\sum Q_m / \sum Q$
$Q / \sum Q_m$	$\frac{\langle Q_{m1} \rangle - \langle Q_{m2} \rangle}{\langle Q_{m1} \rangle + \langle Q_{m2} \rangle}$ $D_{m1 \rightarrow m2}^{-1}$ $\mathbb{A}$ $\mathbb{B}$	$\langle t_m \rangle$

TABLE 5.2 – Features on the edges on the graph. It use the same notation as in table 5.1.  $D_{m1 \rightarrow m2}^{-1}$  is the inverse of the distance between the mesh  $m1$  and the mesh  $m2$ . The features  $\mathbb{A}$  and  $\mathbb{B}$  are detailed in Section 5.1

2315 Since our different nodes do not have the same number of features, they exist in distinct spaces.  
 2316 Traditional graph neural networks only handle homogeneous graphs, where the nodes and edges  
 2317 have the same number of features at each layer. Therefore, the libraries and publicly available  
 2318 algorithms we found were not suited to our needs. As a result, we had to develop and implement a  
 2319 custom message-passing algorithm capable of handling our heterogeneous graph.

## 2320 5.2 Message passing algorithm

2321 The message passing algorithm define the way the GNN will compute and update its graph. As it is  
 2322 detailed in Section 3.2.3, the message-passing algorithm allows each node in the graph to update its  
 2323 features based on information from its neighboring nodes. This update process enables the network  
 2324 to propagate information through the graph, allowing nodes to gradually integrate knowledge about  
 2325 the entire detector. This step is crucial for ensuring that each node can take into account not only its  
 2326 local neighborhood but also the broader context of the event.

2327 As introduced in previous section and in the tables 5.1 and 5.2, our graphs nodes and edges will  
 2328 have different number of features depending on their nature, meaning that we cannot have a single  
 2329 message passing function. We thus need to define a message passing function for each transition  
 2330 inside or outside a family. Using the notation presented in Section 3.2.3:

$$n_i^{k+1} = \phi_u(n_i^k, \square_j \phi_m(n_i^k, n_j^k, e_{ij}^k)); n_j \in \mathcal{N}'_i \quad (5.5)$$

and denoting the mesh nodes  $M$ , the fired nodes  $F$  and the I/O node  $IO$ , we need to define

$$\begin{aligned} & \phi_{u;F \rightarrow M}; \phi_{m;F \rightarrow M} \\ & \phi_{u;M \rightarrow F}; \phi_{m;M \rightarrow F} \\ & \phi_{u;M \rightarrow M}; \phi_{m;M \rightarrow M} \\ & \phi_{u;M \rightarrow IO}; \phi_{m;M \rightarrow IO} \\ & \phi_{u;IO \rightarrow M}; \phi_{m;IO \rightarrow M} \end{aligned}$$

2331 to update the nodes after each layers. Following the illustration in Figure 5.3, for each transition  
 2332 between families or inside a family we need an aggregation, a message and an update function. For  
 2333 the aggregation, we use the sum. We use the same, simple, formalism for every  $\phi_u$ :

$$\phi_u \equiv I_{i'}^{n'} = I_i^n A_{i',e}^i W_n^{e,n'} + I_i^n S_n^{n'} + B^{n'} \quad (5.6)$$

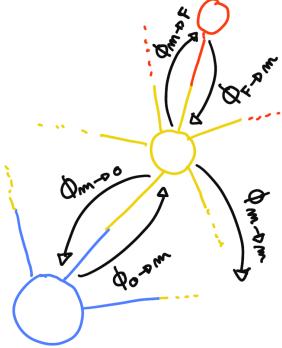


FIGURE 5.3 – Illustration of the different update function needed by our GNN

using the Einstein summation notation. The second order tensor, or matrix,  $I_i^n$  is holding the nodes informations with  $i$  the node index and  $n$  the feature index.  $n$  represent the features of the previous layer and  $n'$  the features of this layer.

$A_{i',e}^i$  is the adjacency tensor, discussed in the previous section, representing the edges between the node  $i'$  and the node  $i$ , each edges holding the features indexed by  $e$ . If the edge does not exist, the features are set to 0. This choice is justified by the linearity of the operation in equation 5.6 : whatever the weights, when multiplied by 0 the results is 0 and the sum result is unchanged.

The learnable parameters are composed of:

- The third order tensor  $W_n^{e,n'}$  which represent the passage from the previous combined feature space between the node and the edge features  $n \otimes e$ , the previous layer, to the current space  $n'$ , this layer.
- The first order tensor  $B^{n'}$  which is a learnable bias on the new features  $n'$ .
- The second order tensor  $S_n^{n'}$ , which can be viewed as a self loop relation where the node update itself based on the previous layer informations, going from the previous space  $n$  to the current space  $n'$ .

If a node have neighbours in different families, the different IAW coming from the different families are summed.

$$I' = \sum_{\mathcal{N}} [I_{\mathcal{N}} AW] + IS + B \quad (5.7)$$

where  $\mathcal{N}$  are the neighbouring family. In our case, dropping the tensor indices and indexing by family for readability, we get

$$I'_F = I_M A_{M \rightarrow F} W_{M \rightarrow F} + I_F S_F + B_F \quad (5.8)$$

$$I'_M = I_F A_{F \rightarrow M} W_{F \rightarrow M} + I_M A_{M \rightarrow M} W_{M \rightarrow M} + I_{IO} A_{IO \rightarrow M} W_{IO \rightarrow M} + I_M S_M + B_M \quad (5.9)$$

$$I'_{IO} = I_M A_{M \rightarrow IO} W_{IO \rightarrow M} + I_{IO} S_{IO} + B_{IO} \quad (5.10)$$

We thus have a  $S$ ,  $W$  and  $B$  for each of the  $\phi_u$  function we defined above. The IAW sum can be seen as the  $\phi_m$  function and  $IS + B$  as the second part of the  $\phi_u$  function. Eq 5.5 gave the generic form of message passing : to update a node  $i$ , one first combines informations from the surrounding nodes and edges and then combine the result ( $\square_j \phi_m$ ) with the current features of node  $i$ . Many practical ways to combine can be tried. In our implementation of message passing (Eq. 5.6 and 5.7) the latter combination is the simple sum of the former (IAW, the equivalent of  $\square_j \phi_m$ ) with a linear combination of the current features of node  $i$  ( $IS + B$ ).

Interestingly, the number of learnable weight in those layer is independent of the number of nodes in each family and depends solely on the number of features on the nodes and the edges.

The expression above only update the node features. We could update the edges, using the results of  $\phi_m$  for example, but for technical simplicity we only update the nodes and keep the edges constant.

Preserving the edges after each layers allow to share the adjacency matrix between all layers, saving memory and computing time.

This operation of message passing is the constituent of our message passing layers, designed in this work as *JWGLayer*, each of them owning their own set of parameter  $W$ ,  $S$  and  $B$ . To those layers, we can adjoin an activation function such as *PReLU*

$$I' = \text{PReLU} \left( \sum_N \left[ I_N A W \right] + I S + B \right) \quad (5.11)$$

### 5.3 Data

The dataset consists of 1M simulated positron events from the JUNO official simulation version J23.0.1-rc8.dc1. This version of the simulation incorporates both the physics of the detector and its electronics, ensuring that the events closely reflect real detector conditions. Importantly, this version includes advanced digitization and trigger modeling, making it suitable for testing the reconstruction capabilities of our GNN model. Those events are uniformly distributed in energy with  $E_k \in [0, 9]$  MeV and distributed in the detector.

All the events are *calib* level, with simulation of the physics, electronics, digitizations and triggers. 900k events will be used for the training, 50k for validation and loss monitoring and 50k for the results analysis in Section 5.7. Each event is between 2k and 12k fired PMTs, resulting in fired nodes being the largest family in our graphs in all circumstances as illustrated in Figure 5.4c.

As expected, by comparing the scale between the Figure 5.4a and 5.4b we see that the LPMT system is predominant in term of informations in our data. The number of PMT hits grow with energy but do not reach 0 for low energy event due to the dark noise contribution which seems to be around 1000 hits per event for the LPMT system (left limit of Figure 5.4a) and around 15 hits per event for the SPMT system (left limit of Figure 5.4b) which is consistent with the results shown in Section 4.1.2.

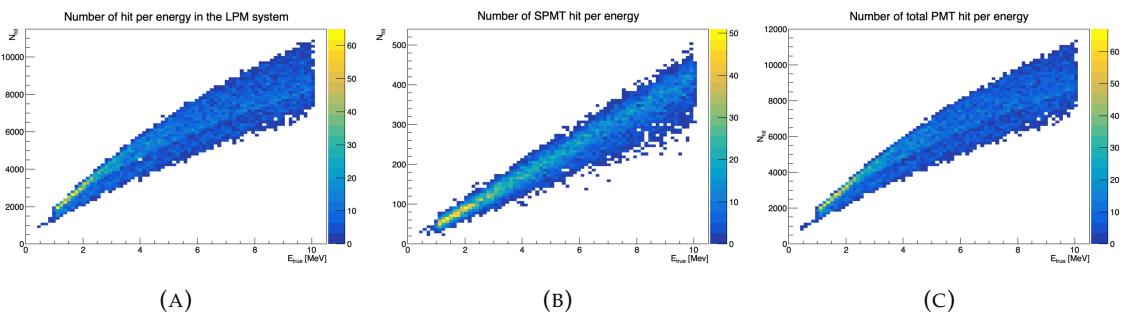


FIGURE 5.4 – Distribution of the number of hits depending on the energy. **On the right:** for the LPMT system. **In the middle :** for the SPMT system. **On the left:** For both system.

The structure seen in the distribution in Figure 5.4a comes from the shape of the number of hits depending on the radius as shown in Figures 5.5a and 5.5b where the number of hit decrease with radius. It is important to understand that this is not representative of the number of PE per event and the decrease in hits over the radius means that the PE are just more concentrated in a smaller number of PMTs.

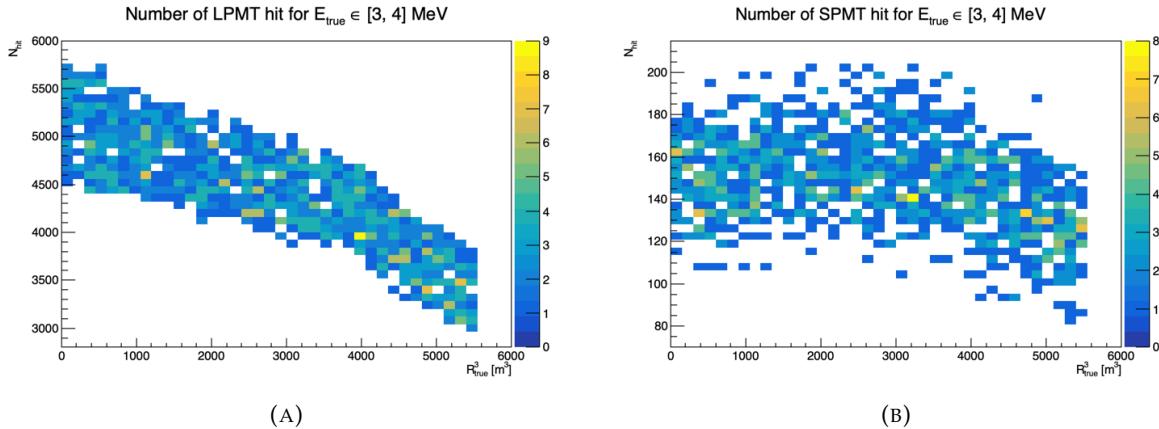


FIGURE 5.5 – Distribution of the number of hits depending on the radius. **On the right:** for the LPMT system. **On the right :** for the SPMT system. To prevent the superposition of structure of different scales we limit ourselves to the energy range  $E_{true} \in [0, 9]$ .

2388 No quality cut is applied here, we rely only on the trigger system. It means that event that would not  
 2389 trigger are not present in the dataset but for events that triggered twice, it happens rarely, the two  
 2390 trigger are considered as two separate event.

## 2391 5.4 Model

2392 In this section, we discuss the different layers that compose the final version of the model. The num-  
 2393 ber of layers, their dimensions, and their arrangement were fine-tuned through multiple iterations.  
 2394 As mentioned earlier, each JWGLayer is defined by the number of features on the nodes and edges  
 2395 of the output graph, assuming it takes as input the graph from the previous layer. For simplicity,  
 2396 when discussing a graph configuration, it will be presented as follow: {  $N_f$ ,  $N_m$ ,  $N_{IO}$ ,  $N_{f \rightarrow m}$ ,  $N_{m \rightarrow m}$ ,  
 2397  $N_{m \rightarrow f}$  } where

- 2398 —  $N_f$  is the number of feature on the fired nodes.
- 2399 —  $N_m$  is the number of features on the mesh nodes.
- 2400 —  $N_{IO}$  is the number of features on the I/O node.
- 2401 —  $N_{f \rightarrow m}$  is the number of features on the edges between the fired and mesh nodes.
- 2402 —  $N_{m \rightarrow m}$  is the number of features on the edges between two mesh nodes.
- 2403 —  $N_{m \rightarrow f}$  is the number of features on the edges between the mesh nodes and the I/O node.

2404 Because we do not change the number of features on the edges, we can simplify the notation to {  $N_f$ ,  
 2405  $N_m$ ,  $N_{IO}$  }. As an example, the input graph configuration, following the tables 5.1 and 5.2 is { 6, 8, 13 },  
 2406 5, 8, 5 } or, without the edge features, { 6, 8, 13 }.

2407 The final version of the model, called JWGV8.4.0 is composed of

- 2408 — An JWGLayer, converting the input graph { 6, 8, 13 } to { 64, 512, 2048 } with a PReLU activation  
 2409 function.
- 2410 — 3 resnet layers, each of them composed of
  - 2411 1. 2 JWG layers with a PReLU activation function. They do not change the dimension of the  
 2412 graph
  - 2413 2. A sum layer that sums the features in the input graph with the one computed from the  
 2414 JWG layers

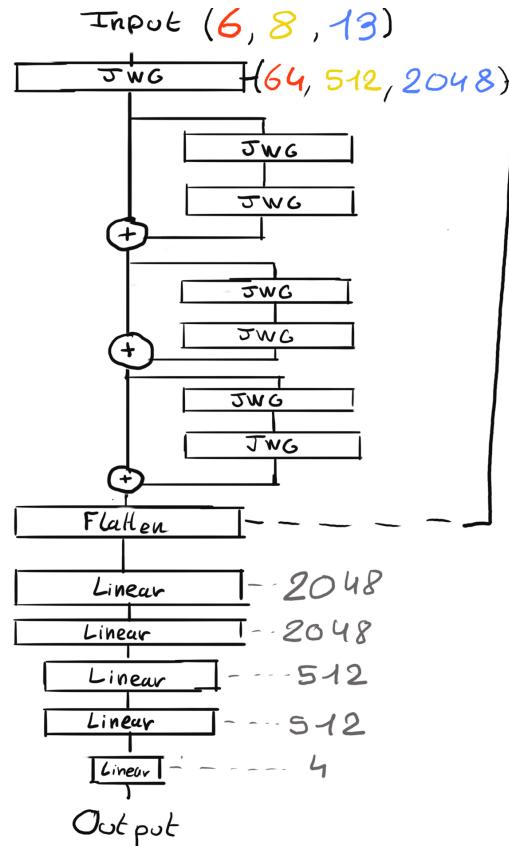


FIGURE 5.6 – Schema of the JWGv8.4.0 architecture, the colored triplet is the graph configuration after each JWG layers

- 2415 — A flatten layer that flatten the features of the I/O and mesh nodes in a vector.
- 2416 — 2 fully connected layers of 2048 neurons with a PReLU activation function.
- 2417 — 2 fully connected layers of 512 neurons with a PReLU activation function.
- 2418 — A final, fully connected layer of 4 neurons acting as the output of the network.

2419 A schematic of the model is presented in Figure 5.6.

2420 We use the Mean Square Error (MSE) for the loss

$$\mathcal{L} = (E_{rec} - E_{dep})^2 + (X_{rec} - X_{true})^2 + (Y_{rec} - Y_{true})^2 + (Z_{rec} - Z_{true})^2 \quad (5.12)$$

2421 as it was the best resulting loss in Chapter 4.

## 2422 5.5 Training

2423 The optimizer used for training is the Adam optimizer (see Section 3.1.3) and default hyperparameters ( $\beta_1 = 0.9$ ,  $\beta_2 = 0.999$  and  $\epsilon = 1e-8$ ) with a learning rate  $\lambda = 1e-8$ . The training last 200 epochs  
2424 of 800 steps. We use a batch size of 32, the largest we can have with 40GB of GPU ram. The learning  
2425 rate is constant during the first 20 epochs then exponentially decrease with a rate of 0.99. We save

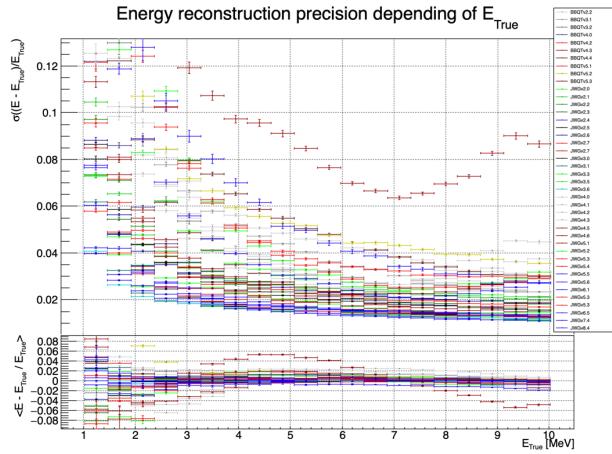


FIGURE 5.7 – Energy reconstruction depending on the true energy for samples of the different versions of the GNN

two set of parameters, the set of parameters the set that yield the lowest validation loss and the set of parameters at the end of the training. The validation is computed over a single batch.

## 5.6 Optimization

The GNN model presented in previous sections is the result of a long work of optimization. Indeed, the innovative architecture we propose left us with an infinity of possible configurations with no guidance from prior works in literature nor in JUNO.

In the end, more than 60 different configurations have been tested. This effort is illustrated on Figure 5.7<sup>1</sup>, where the 40 configurations are compared in their ability to reconstruct the positron energy. Although all configurations share the fundamental principles we base our innovative architecture on (three different kinds of nodes and edges, usage of raw level features on some of them, usage of higher level data on others, division of JUNO’s surface into regional pixels to form mesh nodes, the very large number of edges connected to each mesh node, etc.), performances can vary a lot between our first attempts (far beyond any acceptable energy resolution, and not even on this figure) and recent ones. Therefore: the precise way to choose hyperparameters mattered a lot, regardless of the relevance of the global architectural principles.

The spectacular improvement between early and later configurations also explains the length of this process : for long we hoped we would finally reach the classical performance, and it was tempting to test yet another configuration.

### 5.6.1 Software optimization

A substantial effort was devoted to the data processing workflow. Transforming JUNO simulation outputs into graphs is a computationally expensive task. Furthermore, due to the ever-changing nature of the graph dimensions and features during optimization, preprocessing JUNO’s files by precalculating the graphs and then reading them from files was not viable, as it would require a large amount of disk space to store events for each version of the graph.

1. Note that this figure was prepared on idealized data with no dark noise and perfect hit time determination.

2451 Therefore, the software does not rely on preprocessed data and instead computes the observables,  
 2452 adjacency matrix, etc., during training. This data processing is performed in parallel on the CPU.  
 2453 The raw data comes from ROOT files produced by the collaboration software, and the Event Data  
 2454 Model (EDM), used internally by the collaboration [128], had to be interfaced with our software,  
 2455 an interface that had to be maintained as the collaboration's software evolved. For the harmonic  
 2456 power calculation, we migrated from the Healpix library to Ducc0 [129] for more precise control  
 2457 over multithreading.

### 2458 5.6.2 Hyperparameters optimization

2459 The first kind of hyper-parameters that received a lot of effort concern the network's detailed architecture:

- 2461 — Message passing layers where originally not JWG layers, we started by using small FCDNN in  
 2462 place of  $\phi_u$  and  $\phi_m$ . Due to low performances and memory consumption issues, we pivoted to  
 2463 the message passing algorithm presented in Section 5.2.
- 2464 — The ResNet architecture was brought after issue with the gradient vanishing.
- 2465 — The number of layers was varied between 5 and 12.
- 2466 — The number of node features after each given message passing layer (64, 512, 2048 in the final  
 2467 version) was varied.
- 2468 — The Final FCDNN after the message passing layers is not present in all versions.
- 2469 — At some point, the PReLU activation function replaced the ReLU function.

2470  
 2471 For some of them, software work was necessary. In any case, each configuration required a training  
 2472 of about 90h. Adding the analysis time necessary to the verification of its performance and the  
 2473 comparison with other versions, one understands the number of tests had to be limited.

2474 Other hyperparameters were also tested :

- 2475 — The higher level variables described in Section 5.1 (powers of various spherical harmonics,  $\mathbb{A}$ ,  
 2476  $\mathbb{A}$ ,  $(Q_{m1} - Q_{m2})/(Q_{m1} + Q_{m2})$  ) were added progressively. Notice that our choice to focus  
 2477 our search on this kind of variables is also due to the fact that JWGLayer involves linear  
 2478 operations. It is therefore difficult for such a network to propose variables of this kind among  
 2479 the node features learned layers after layers (i.e. it's difficult for the network to understand  
 2480 these variables are important, or only after many layers).
- 2481 — Time allocated to training, the Learning Rate, the size of batches, etc.
- 2482 — The number of pixels (ie of mesh nodes) was varied between 192 and 768.
- 2483 — Several definitions loss functions where tried. In particular, we tried some focussed only on  
 2484 the E resolution, only on the vertex resolution (R) or trying to optimize both.

2485

2486 To make a long story short, each new configuration was the result of our reflections after having  
 2487 analysed the previous configurations, or after having thought over again about JUNO's detailed  
 2488 response to energy deposits – seeking for variables that could help the GNN.

2489 Another, quite common, approach was in principle possible : a random search. However, due to the  
 2490 extensive training time, up to 90h per training, the heavy memory consumption of the models that  
 2491 would often exceed the 20GB limit of the V100, this approach was not realistic in our case, though we  
 2492 were able to extend the memory limit to 40GB thanks to a local A100 GPU card available at Subatech.

## 2493 5.7 performance of the final version

2494 The reconstruction performance of "JWGv8.4" are presented in Figures 5.8, 5.9, 5.10 and compared  
 2495 to the "Omilrec" algorithm, the official IBD reconstruction algorithm in JUNO. Omilrec is based on  
 2496 the QTMLE reconstruction method that was presented in Section 3.3.

2497 This comparison required to use a consistent definition of  $E_{true}$ . This is not trivial since at JUNO,  
 2498 ML method reconstruct the true energy deposited by the positron+annihilation gammas (that's the  
 2499 target implemented in the loss function), while Omilrec, which is based on probabilities to observe a  
 2500 given number of PE in a given PMT, reconstruct the "visible energy". It reflects the total number of  
 2501 radiated and detectable scintillation or Cherenkov photons (and is subject to non linear effects like  
 2502 quenching).

2503 The conversion we use to obtain comparable  $E_{true}$  is explained in Appendix C.

2504 On Figures 5.8 to 5.10, we notice that the best GNN does not match the performance of the OMILREC  
 2505 algorithm. Generically, Energy resolution is 50% worse, while the resolution on R is three times  
 2506 worse. Reconstruction biases are not better either with the GNN. We have tried to understand the  
 2507 origin of this limited performance.

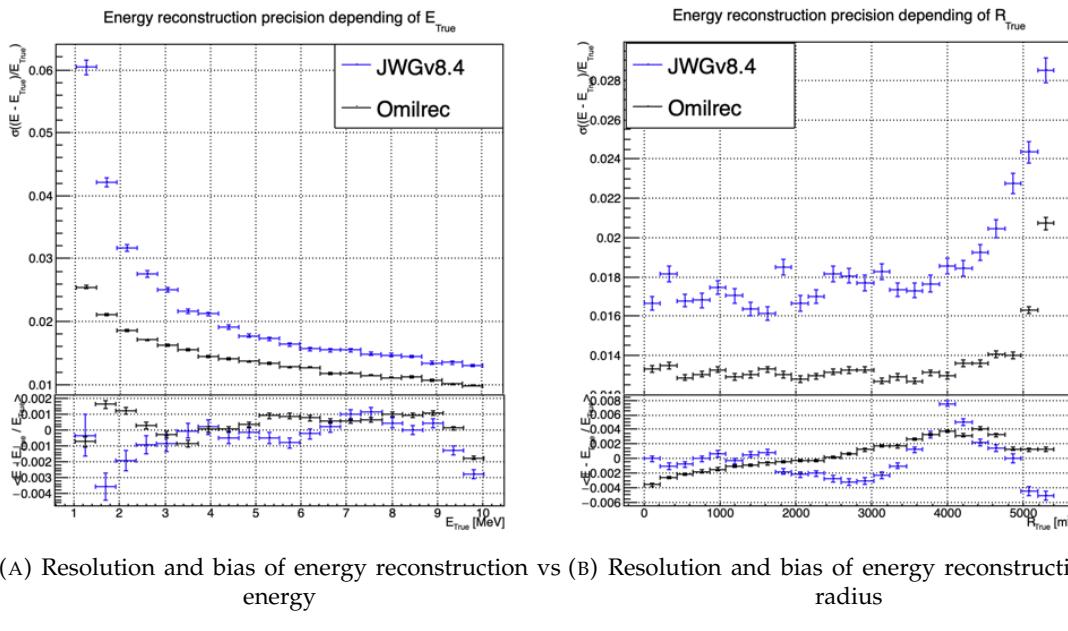


FIGURE 5.8 – Reconstruction performance of the Omilrec algorithm based on QTMLE presented in Section 3.3, JWGv8.4 presented in this chapter. The top part of each plot is the resolution and the bottom part is the bias.

2508 The first action that can be carried out in this direction was to determine if some information used  
 2509 by OMILREC was not used properly by JWGv8.4. For that purpose, we used again the approach  
 2510 presented in Chapter 4 (Sec 4.3.2 and annex A ) to combine JWGv8.4 and OMILREC. We observe on  
 2511 Figures 5.11 and 5.12 that this combination brings no sizeable improvement of the best of the two  
 2512 combined methods. The combination remains very close to OMILREC alone. This is an indication  
 2513 that JWGv8.4 does not use informations that would be overlooked by OMILREC, and that on the  
 2514 contrary, that's JWGv8.4 that fails to use properly important informations.

2515 The problem described above could be inherent to our GNN's original architecture. Discussions with  
 2516 JUNO's colleagues when these results were presented at the collaboration pointed to the role of PMT  
 2517 time information ( $t$ , in the  $(Q, t)$  pairs we use as our algorithm input features). The thousands of

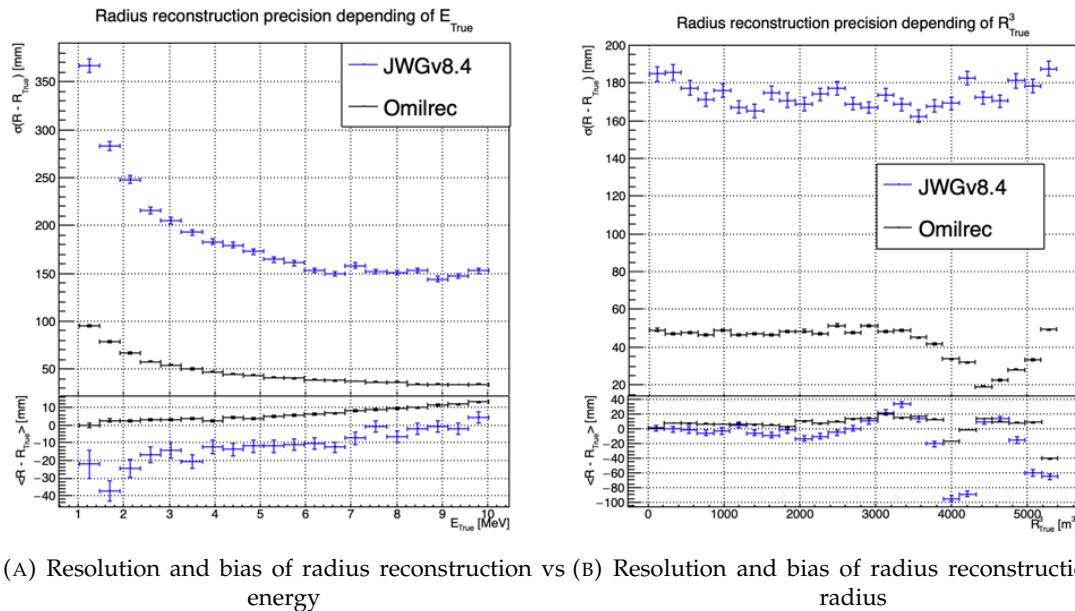


FIGURE 5.9 – Reconstruction performance of the Omilrec algorithm based on QTMLR presented in Section 3.3, JWGv8.4 presented in this chapter. The top part of each plot is the resolution and the bottom part is the bias.

values found in the *fired* nodes might not be aggregated well enough when transmitted to the mesh nodes, causing a loss in the redundancy of this important information.

We tested this idea in several manners, described below.

### Finer granularity

We tried to recover some redundancy by increasing the number of mesh nodes from 198 to 768. The improvement we observed was small, and did not allow to get close to OMILREC's performance.

To explore further in this direction, we would ideally try 3072 pixels (the next HEALPIX rank). However, this is not possible for our GNN due to hardware limitations, mainly the available GPU memory. Instead, we discussed the problem with Gilles Grasseau, calculus research engineer with whom we collaborate on the subject of ML reliability (see Chapter 6). In the framework of this activity, Gilles needs to develop reconstruction algorithms to be "attacked" by a prototype Adversarial NN. One of them is a pseudo-spherical CNN using oriented filters, called HCNN.

To produce its input image, this algorithms split the Sphere into 3072 pixels. Each channel of this image is an aggregation of the  $(Q, t)$  values found in all the PMTs. The charge are summed and the lowest time is kept. The performance of this algorithm can be seen on Figures 5.13 and 5.14, compared to OMILREC. With 3072 pixels, the performance of HCNN does not match that of OMILREC, but is closer to it than our GNN. The granularity of the pixels, and the way to summarize the individual PMTs information when going from 17000 LPMTs to only 3072 pixels indeed seems to play a role.

This is consistent with the results obtained by the first GNN tried at JUNO on reactor neutrinos (already described in Section 3.3.3). It used 3072 pixels, and also obtained an uncompetitive R reconstruction.

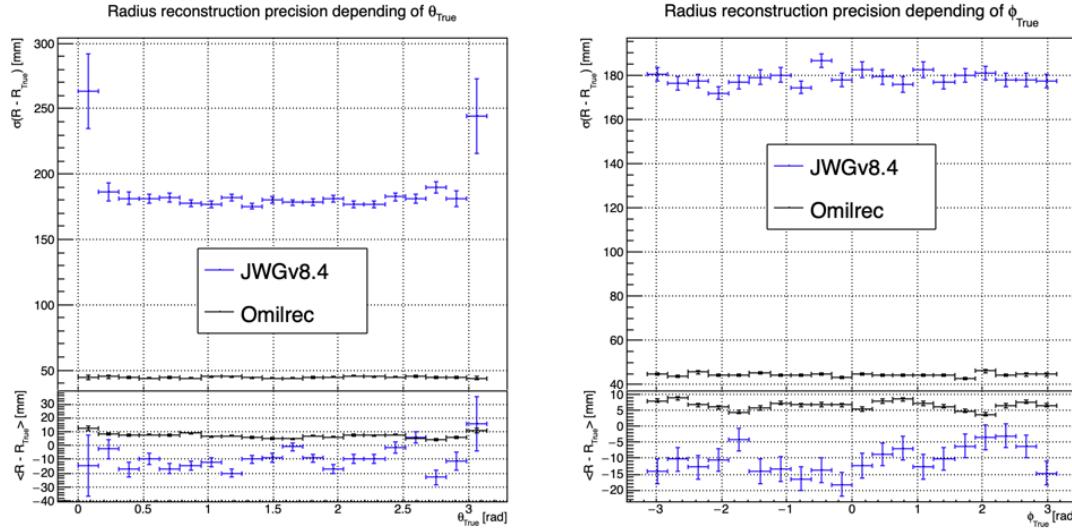


FIGURE 5.10 – Reconstruction performance of the Omilrec algorithm based on QTMLE presented in Section 3.3, JWGv8.4 presented in this chapter. The top part of each plot is the resolution and the bottom part is the bias.

## 2540 Information reduction, from fired to Meshes

2541 The problem described above is somehow classical. ML algorithms, ideally, would start from the  
 2542 full information present in the detector, and learn to reduce it optimally.

2543 In cases where only 3072 pixels can be used instead of the complete information from 17000 PMTs,  
 2544 one needs to understand how to combine the individual from the 5 or 6 PMT found in each pixel into  
 2545 pixel-level features, without loosing important information.

2546 In the case of our GNN, we hoped that by connecting each mesh node to its corresponding 5 or 6  
 2547 fired nodes, we could keep the full information. In reality, it seems that the message passing between  
 2548 fired and mesh does not work efficiently. When nodes are updated by the first (may be also by the  
 2549 subsequent) layer, the new mesh features might be dominated by the original features in the second  
 2550 column of tables 5.1, themselves a simple version of aggregation. Layer after layer, we might be  
 2551 limited to that level of time information, lacking time redundancy.

2552 We have verified this by testing version of the GNN in which the link between fired and mesh was  
 2553 cut, or in which no time info was included among the fired nodes features. It had only a small effect  
 2554 which seems to confirm a problem in the way the full information, from all the individual PMTs, is  
 2555 used by our GNN.

## 2556 Possible improvements

2557 It appears that the network is unable to aggregate the timing information correctly. While this could  
 2558 be addressed by using a finer segmentation, with more mesh nodes, improvements might also arise  
 2559 from refining the message-passing algorithm. The algorithm presented in this thesis is still quite  
 2560 basic, relying on a simple linear combination of features. We have seen through examples in CNNs,  
 2561 GNNs, and other architectures, both in research and industry, that specializing the network – for  
 2562 instance, by incorporating convolutional filters – can lead to improvements that were previously

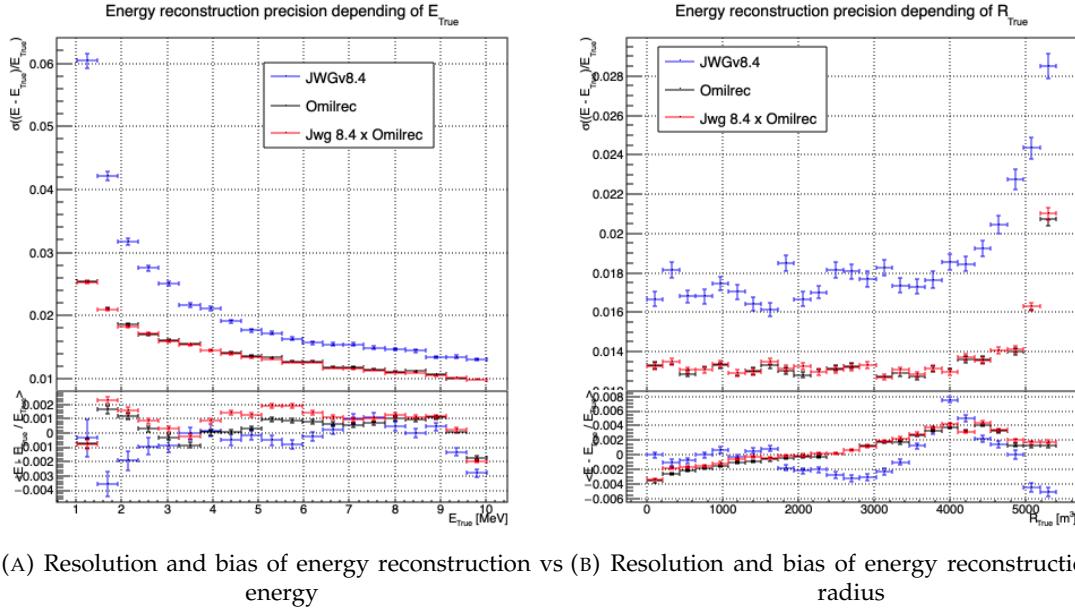


FIGURE 5.11 – Reconstruction performance of the Omilrec algorithm, JWGV8.4 and the combination between the two using the optimal variance estimator presented in annex A.2. The top part of each plot is the resolution and the bottom part is the bias.

unattainable with simpler FCDNNs. Applying this approach to the message-passing algorithm, by utilizing a GNN with a more advanced message-passing, could yield better results.

We could investigate alternative aggregation strategies, for example, by weighting the timing information more significantly during the message-passing phase. Additionally, testing a non-linear combination of features from fired to mesh nodes could help preserve more granular information. Another potential improvement would be to introduce attention mechanisms that dynamically assign more importance to relevant features in the fired nodes

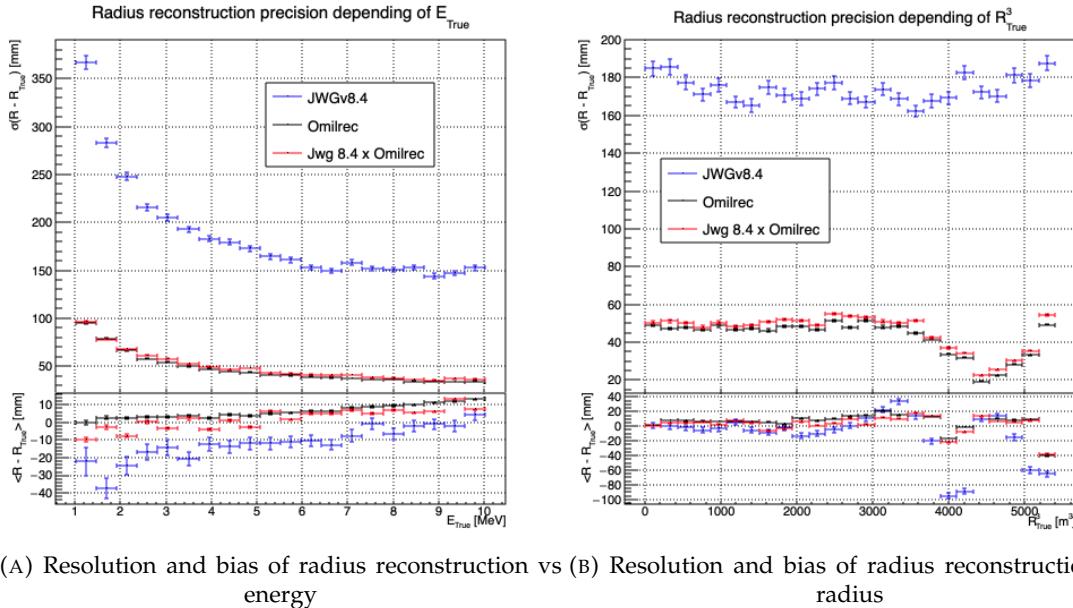
Regarding the timing information, we provided high-level features, assuming this would assist the neural network in converging to the solution. However, by offering such information upfront, the GNN might be taking the “easy” path, settling for a local and broader minimum, rather than extracting the features that could lead to better performance.

If there are difficulties in transferring information between the fired and mesh nodes, it may stem from the way we connected the fired nodes to the mesh nodes. By linking the fired nodes within the same mesh, or even connecting the fired nodes of neighboring mesh nodes, the GNN might be able to construct more meaningful information.

Finally, by providing directly the PMT waveform to the GNN, in the fired nodes, we could search for even finer precision and results. An idea would be to specialise the message function  $\phi_{m;F \rightarrow M}$  to be a 1D convolutional layer over the waveform. The resulting channels would be fed to the mesh nodes for their updates.

## 5.8 Conclusion

To achieve its scientific goals, JUNO requires a precise and well-understood reconstruction, as it needs an energy resolution of 3% at 1 MeV. Even small, unaccounted biases could make it impossible to determine the mass ordering, as explored in Chapter 7. A likelihood-based algorithm, designed to



(A) Resolution and bias of radius reconstruction vs (B) Resolution and bias of radius reconstruction vs  
energy radius

FIGURE 5.12 – Reconstruction performance of the Omilrec algorithm, JWGV8.4 and the combination between the two using the optimal variance estimator presented in annex A.2. The top part of each plot is the resolution and the bottom part is the bias.

meet JUNO’s requirements and referred to as the classical algorithm, was developed and is detailed in Section 3.3.

Machine learning algorithms were developed to challenge this classical approach, and they are presented in Section 3.3.3. Although they achieve the precision of the classical algorithm, they do not offer significant improvements. The GNN previously developed is a convolutional GNN where nodes correspond to pixels, connected to their neighbors based on the Healpix [127] segmentation, with the  $(Q, t)$  information aggregated onto these pixels.

In this chapter, we introduce a novel and innovative architecture. In addition to the pixel segmentation represented by mesh nodes, we incorporate rawer information by directly representing the fired PMTs as nodes. We also fully connect the mesh nodes to each other, hoping to facilitate the transfer of information. Finally, we introduce a global node that holds global information about the detector.

These three types, or families, of nodes do not have the same number of features, resulting in a heterogeneous graph. Publicly available algorithms for graph processing are designed for homogeneous graphs, so we had to develop a custom algorithm adapted to heterogeneous graphs.

This GNN required significant technical development, but the results are not at the level of the classical algorithm. The tests we conducted suggest that the problem may lie in the aggregation of raw information from the fired nodes onto the mesh nodes, as removing the fired nodes does not degrade the results. Additionally, due to technical constraints, we had to reduce the number of pixels compared to the previous GNN. Other algorithms we developed, which use a higher pixel resolution, outperform this architecture, reinforcing our suspicion that the aggregation is the root of the issue.

The precision required for JUNO’s scientific objectives, particularly in determining mass ordering, imposes stringent constraints on reconstruction algorithms. Small biases or errors in energy resolution could significantly affect the experiment’s outcomes. Future improvements may involve refining the message-passing algorithm, incorporating additional detector-specific features, and experimenting with more advanced architectures such as attention-based GNNs to further reduce reconstruction

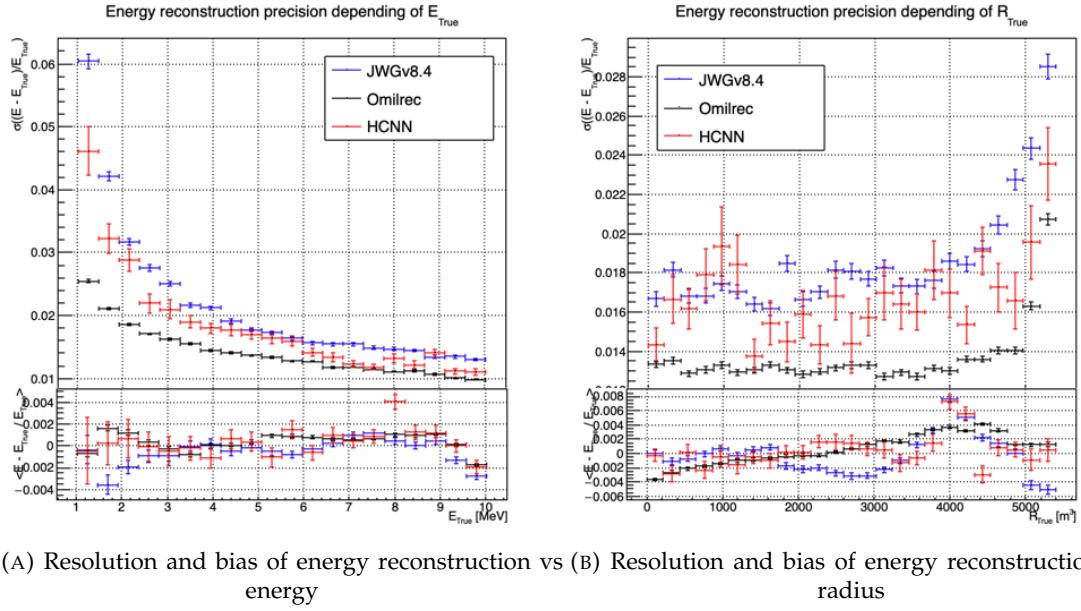


FIGURE 5.13 – Reconstruction performance of the Omilrec algorithm based on QTMLE presented in Section 3.3, JWGv8.4 presented in this chapter and the HCNN algorithm. The top part of each plot is the resolution and the bottom part is the bias.

2612 errors.

2613 Perhaps by incorporating rawer information, such as the waveform, refining the message-passing  
 2614 algorithm, or adjusting the features on the different nodes, we could match the precision of the  
 2615 classical algorithm. However, it is also possible that deeper, more radical changes are needed to  
 2616 become competitive.

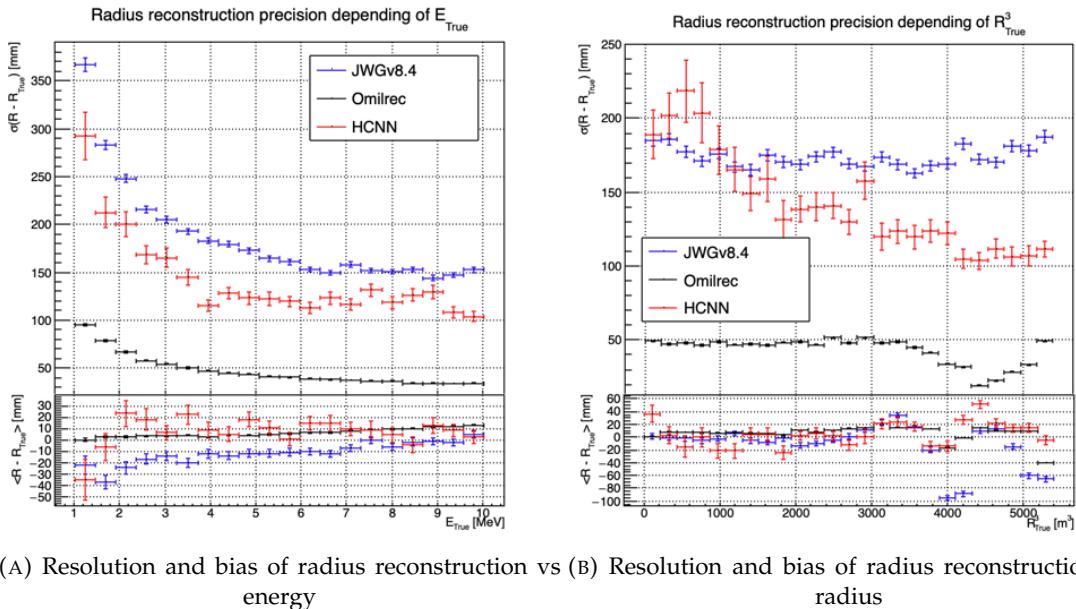


FIGURE 5.14 – Reconstruction performance of the Omilrec algorithm based on QTMLE presented in Section 3.3, JWGV8.4 presented in this chapter and the HCNN algorithm. The top part of each plot is the resolution and the bottom part is the bias.

<sup>2617</sup> **Chapter 6**

<sup>2618</sup> **Reliability of machine learning  
methods**  
<sup>2619</sup>

<sup>2620</sup> “*Psychohistory was the quintessence of sociology; it was the science of  
human behavior reduced to mathematical equations. The individual  
human being is unpredictable, but the reactions of human mobs,  
Seldon found, could be treated statistically*”

*Isaac Asimov, Second Foundation*

<sup>2621</sup> **Contents**

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<sup>2629</sup> <b>6.3 Conclusion and prospects . . . . .</b>	<b>127</b>

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<sup>2630</sup> As explained in previous chapters, JUNO is a precision experiment where a very precise understanding of the reconstruction effects is crucial. <sup>2631</sup> JUNO is a high-precision experiment, and any <sup>2632</sup> discrepancies in the understanding of reconstruction effects can significantly impact the neutrino <sup>2633</sup> mass ordering (NMO) determination. Particular attention must be paid to potential biases arising <sup>2634</sup> from energy scale mis-calibration, charge non-linearity, and differences between real and simulated <sup>2635</sup> detector responses, which could skew the oscillation results.

<sup>2636</sup> While the liquid scintillator technology is well known, this is the first time it is deployed to such scale, <sup>2637</sup> and for such precision. This novelty bring makes this task difficult : a lot of effects should ideally <sup>2638</sup> be understood better than in previous experiments. We already know that a bad knowledge of the <sup>2639</sup> energy scale can have consequences as serious as excluding the wrong NMO. It must be known at <sup>2640</sup> the 1% level to reach the desired sensitivity. The present chapter is motivated by a specific question <sup>2641</sup> : even very small differences between reconstruction effects in real data and in the models used in <sup>2642</sup> oscillation fits could alter this sensitivity or bias the results. Reconstruction algorithms developed <sup>2643</sup> in JUNO, in particular those based on ML, try to use the information present in the detector as <sup>2644</sup> exhaustively as possible. This might apply even more to future algorithms. There is here a risk that <sup>2645</sup> some of the used information is not similar in real data and in the detector’s model, leading to the <sup>2646</sup> problem above mentioned. We worked on two ways to address this problem.

<sup>2647</sup> We think the simpler way to study this reliability issue is to compare algorithm with each other. <sup>2648</sup> Differences between various reconstructed spectra of the same sample provide an envelope to evaluate <sup>2649</sup> the scale of a potential problem. Comparisons on the event per event basis allows to show that 2 <sup>2650</sup>

algorithms do not use the same information, and is a first step to characterize this difference. We already shown a large variety of reconstruction algorithms, OMILREC for LPMT reconstruction in Section 3.3, numerous machine learning algorithms in Section 3.3.3 and our own work in Chapters 4 and 5. Those algorithms were compared to each others based on their performance as in [114] but we are the first that looked into the correlation between the reconstruction. The combinations of algorithms shown in Section 4.3.2 show that some information elude the algorithms. To efficiently compare algorithms between each other, they need to be publicly available to the collaboration to studies their differences event by event. Nothing of that kind was possible in JUNO up to now. I imported in JUNO's official software some tools necessary to ML algorithms, and I implemented a first algorithms : a for energy reconstruction, named BDTE which was developed by Gavrikov et al. [115], another JUNO's research team. We paved the way for this movement to continue, so that all ML algorithms be available to any JUNO analyst. The details of this implementation and its combination with OMILREC are presented in Section 6.1.

The other way we have explored to study reliability is to challenge reconstruction algorithms with physically plausible perturbations in the PMT charge and time information. More specifically: these perturbations embody differences between the real detector and its model, and we want to identify perturbation patterns which would be too subtle to be detected via data/MC comparisons in calibration data or in control samples from physics data, but which would still be able to alter the oscillation analysis result. We could try to design such perturbations "by hand", based on our knowledge of JUNO. However, with as much as 17600 LPMT, finding subtle perturbations that affecting undefined combinations of PMT would be an endless process.

We propose leveraging machine learning by developing an Adversarial Neural Network (ANN) to introduce perturbations reflecting discrepancies between the real detector and its model. However, one challenge with ANN-based approaches is ensuring that the generated perturbations remain physically plausible and are not overly sensitive to random noise or edge effects in the detector.

In Section 6.2, I describe the method behind the algorithm. In Section 6.2.1 I detail the architecture of our algorithm. The training and the results of our method are presented in Section 6.2.5 . Finally, in section 6.3, I conclude and discuss the prospects and possible improvements to bring to this work.

## 6.1 First implementation of ML methods in JUNO's software

To study the reliability of reconstruction algorithms it's necessary to be able to compare their reconstruction performance event by event. To ease the process, it is important that they are publicly available. JUNO's common software, discussed in Section 2.6, is based on the SNiPER framework [86] which allows the packaging of the different steps of JUNO's analysis, from Monte Carlo (MC) data generation to event reconstruction, including the propagation and interactions of the particles in the LS, the emission and propagation of the scintillation light, the simulation of the PMTs' waveform reconstruction, electronic effects and the trigger system.

This framework is modular, with each module being a C++ class bound in Python. The execution of successive algorithms is orchestrated via Python scripts.

We could have implemented the algorithms presented in Chapters 4 and 5, but since these are themselves not trivial, we chose to start with a simpler ML algorithm that presents similar energy reconstruction performances as OMILREC: a Boosted Decision Tree (BDT) for energy reconstruction developed by Gavrikov Arsenii et al. [115]. This BDT, named BDTE, is based on an aggregated features approach where instead of providing a ML algorithm with low level information, namely the full list of  $(Q, t)$  in LPMTs, a set of higher-level variables is designed based on physicist's common knowledge and then fed to the BDT. The list of the aggregated features used by the BDT is presented in Table 6.1. These higher-order variables are extracted from the charge  $Q$  and hit time  $t$  distribution. It also depends on two straightforward interaction vertex estimators.

2701 The first one is the charge barycenter

$$\vec{r}_{cc} = \frac{\sum_i \vec{r}_{PMT,i} Q_i}{\sum_i Q_i} \quad (6.1)$$

2702 where  $i$  index the fired PMT,  $\vec{r}_{PMT,i}$  is the position vector of the  $i$ th PMT and  $Q_i$  is the charge it  
2703 collected.

2704 The second estimator is the hit time barycenter

$$\vec{r}_{ht} = \frac{1}{\sum_i \frac{1}{t_i + c}} \sum_i \frac{\vec{r}_{PMT,i}}{t_i + c} \quad (6.2)$$

2705 where  $t_i$  is the time of collection of the  $i$ th PMT and  $c = 50$  ns a constant to prevent divergence when  
2706  $t_i$  is 0.

Feature	Description
AccumCharge	Sum of the charge collected by every LPMT
$R_{ht}$	Radius reconstructed by the hit time barycenter
$z_{cc}$	$z$ component of the vertex reconstructed by the charge barycenter
$\sigma_{PE}$	Standard deviation of the distribution of collected PE per PMTs
$N_{PMT}$	Number of fired PMTs
$htKurtosis$	Kurtosis of the hit time distribution
$ht_{25\%}-20\%$	Difference between the 25% and 20% percentiles of the hit time distribution
$R_{cc}$	Radius reconstructed by the center of charge barycenter
$ht_{5\%}-2\%$	Difference between the 5% and 2% percentiles of the hit time distribution
$\langle PE \rangle$	Mean number of PE collected per PMTs
$J_{ht}$	Jacobian of the hit time distribution
$\phi_{cc}$	$\phi$ component in spherical coordinate of the charge barycenter
$ht_{35\%}-30\%$	Difference between the 25% and 20% percentiles of the hit time distribution
$ht_{20\%}-15\%$	Difference between the 20% and 15% percentiles of the hit time distribution
$PE_{35\%}$	Value of the 35% percentile of the charge distribution
$ht_{30\%}-25\%$	Difference between the 30% and 25% percentiles of the hit time distribution

TABLE 6.1 – Summary of the aggregated features used by the BDT to reconstruct the IBD energy. The charge barycenter and hit time barycenter vertex estimators are detailed in Eq. 6.1 and 6.2 respectively

2707 The performance of this BDT, as published by Gavrikov Arsenii et. al, is reported in Figure 6.2a. This  
2708 BDT is developed in Python using the XGBoost [130] library and originally consisted of a collection  
2709 of Python scripts for the training and the evaluation.

2710 As stated before, JUNO software is composed of C++ modules orchestrated through Python scripts.  
2711 The technical challenge was to extract the data from the internal representation of the event in JUNO  
2712 software, the Event Data Model (EDM), into a comprehensible format for Python. This task, which  
2713 was previously done via data pre-processing by Python scripts, had to be internalized within the  
2714 software. The computation of the aggregated features was migrated from the Python scripts into  
2715 C++ modules. The final step was to fetch the reconstruction results of the algorithm into the C++  
2716 framework to save the results in the EDM. Some of the Python libraries were missing, notably  
2717 XGBoost. A request to the collaboration was issued for the packaging of these libraries with the  
2718 common software. As a workaround, the documentation of the algorithm contains the procedure to  
2719 locally install the missing libraries.

2720 We validated the consistency of the aggregated features between the original Python implementation  
2721 and the JUNO software by comparing 1,000 events with the help of Arsenii. For the majority of the  
2722 features, the relative difference between his and ours was either 0 or of the order of  $10^{-15}$ , with the

exception of three features:  $R_{cc}$ ,  $R_{ht}$ , and  $z_{cc}$ . For these three features, the relative difference is about  $10^{-6}$ , which, while small, is still surprisingly high for numerical computation. The distributions of the relative differences for these features are presented in Figure 6.1.

We investigated the source of these discrepancies. The difference in computation environments – Python using Numpy [131] and C++ using the standard library in our case – is most likely the cause. Since the discrepancies arise from the computation of the barycenter in Eq. 6.1 and 6.2, they may result from differences in compiler optimization during the weighted sum calculation. We consider that these differences are still small enough that the performance of the BDT is unaffected.

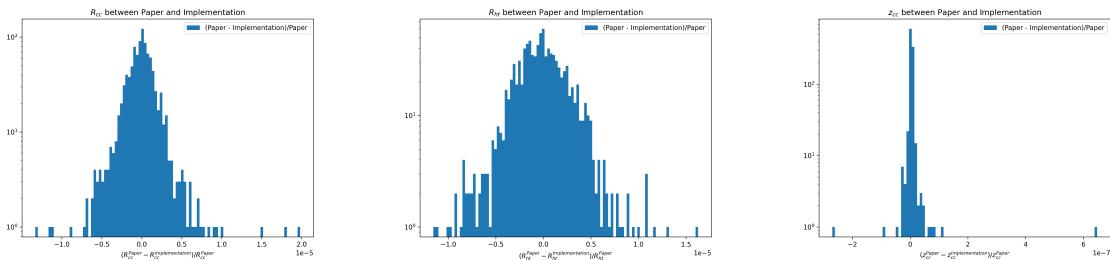


FIGURE 6.1 – Relative difference between the features computed by Gavrikov et. al (superscripted Paper) and our implementation (superscripted Implementation)

The performance of our implementation of BDTE compared to the results presented in [115] are presented in figure 6.2b.

- At 1 MeV, the relative resolution is reported by the publication is just below 3%/MeV. Our implementation show a relative resolution of 2.8%. The relative bias is reported at -0.1%, same as for our implementation.
- At 4 MeV, the reported relative resolution is 1.5%, our implementation show 1.45%. The relative bias is about 0.05% in both results.
- At 10 MeV, our implementation reconstruct the energy with a resolution of 1% whereas the publication report a resolution a bit greater than 1%. They report a positive relative bias of 0.05% while we see a negative 0.1%. This difference might come from the fact that Arsenii provided us an updated version of the BDT since the publication of [115].

The performance are considered compatibles.

The reconstruction using BDTE was implemented in JUNO’s common software but Gavrikov et al. also detail the training and hyper-optimization. JUNO Monte Carlo is likely to evolve during the construction phase and will be further adjusted using calibration. The implementation of those procedures, the training and optimization, will be required as BDTE re-training and re-optimisation will be required with each JUNO software update.

Figure 6.2b shows that the resolution of BDTE is very close to OMILREC. We measured the correlation between their reconstructions, focusing on the residuals with respect to the common true deposited energy:

$$\text{Corr}(E_{BDTE} - E_{dep}, E_{OMILREC} - E_{dep}) \quad (6.3)$$

If the correlation is small enough, it indicates that these two reconstruction algorithms do not use the same information. As a corollary, it indicates that these algorithms can in principle be improved. The correlation between errors for different energy and event radius in the detector is presented in Figure 6.3. We see that for the vast majority of the  $(R^3, E)$  phase space, the correlation is  $> 0.995$ , down to  $\sim 0.98$  in the  $R \approx 9$  m and  $R > 17$  m regions. Such high correlations indicates that these algorithms are very close to using the same information. No difference can be found here, that could be used to improve them. Maybe the situation will be different when other ML algorithms

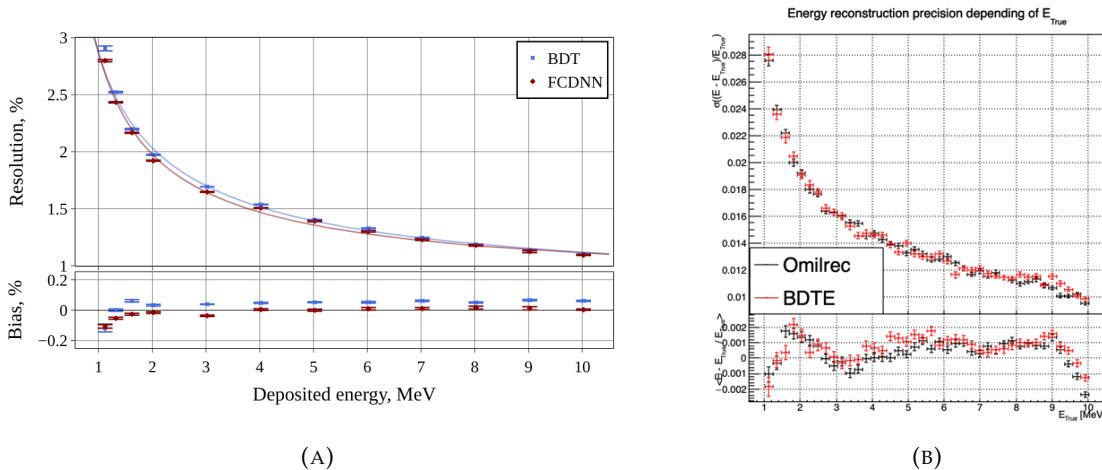


FIGURE 6.2 – Resolution of BDTE **On the left:** as reported by Gavrikov Arsenii et. al in [115], **On the right:** once implemented in JUNO common software. On the right plot is also reported the reconstruction performance of the OMILREC algorithm. The OMILREC algorithm  $E_{\text{vis}}$  has been corrected to  $E_{\text{dep}}$  following the procedure detailed in Annex C.

are implemented in JUNO’s software and the same exercise carried out. The fact that BDTE and OMILREC are so correlated, and their performance so similar, is interesting : it suggests that using the full  $(Q, t)$  list as inputs does not allow major improvements. This is in line with some of the conclusions we expressed in previous chapters : to improve JUNO’s reconstruction by starting from low level variables might requires to use rawer variables than  $(Q, t)$ , like the full waveforms.

## 6.2 Adversarial method

As introduced at the beginning of the chapter, JUNO needs a very good understanding of the biases and effects affecting its reconstruction, and small discrepancies between the real detector and its model could be an issue, in particular when ML algorithms are used. Calibration data will be used to study reconstruction effects and to tune the simulation so that the detector model matches as well as possible the real one. JUNO relies on multiple sources that can be deployed at various positions in the detector. The calibration strategy is already discussed in Section 2.4 and shows calibration sources of gammas, neutrons, and positrons (Table 2.3), with the catch that the positrons will annihilate inside the encapsulation and only the two 511 keV gammas will deposit energy in the LS.

None of the calibration sources considered are positron events. While electrons and positron events should be pretty similar in their interaction with the electronic cloud of the LS atoms, electron events are missing the two annihilation gammas. The topology of the event is therefore not the same: electron events display a single interaction site, up to a few cm long, where a few MeV is deposited; positron events display a similar main site, accompanied with several other low energy ( $< 300$  keV) sites, typically spread over more than 20 cm, due to the Compton interactions of the annihilation gammas. Other differences appear due to the fraction of the positrons that will form a positronium : it causes a delay of a few nanoseconds between energy deposition and the positronium annihilation, to be compared to the PMT transit time spread between 3 and 6 ns, depending on the PMT type [132–134]. Therefore, subtle effects might be present in positron events that the analysis of electron events samples cannot capture. Moreover, not all positions can be reached by calibration sources (effects affecting events close to the border of the detector can’t be studied perfectly then), and calibration run are punctual in time, and therefore can’t witness finely of time evolutions in the

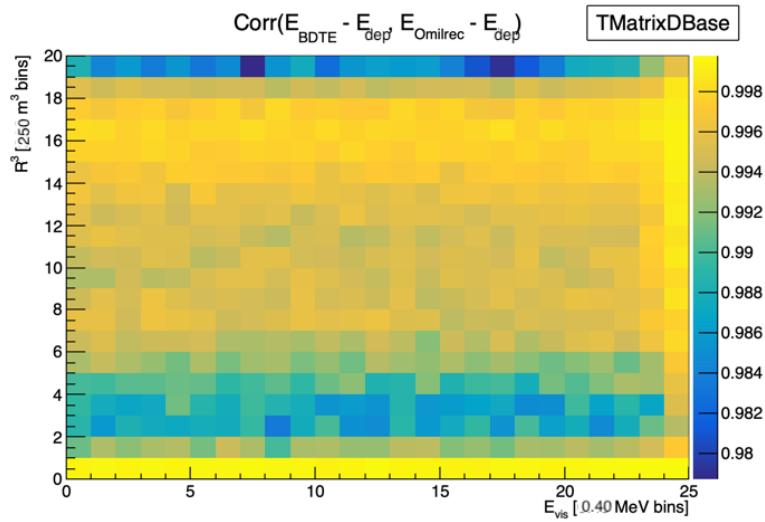


FIGURE 6.3 – Correlation between the errors in energy reconstruction between BDTE and OMILREC (Eq. 6.3). The correlation is computed in  $R^3$  bins of  $216 \text{ m}^3$  between 0 and  $5000 \text{ m}^3$ , 0 and 17 m in y axis, and in 0.40 MeV bins between 1.022 and 10.022 MeV of deposited energy.

2785 detector.

2786 The two last issues presented above do not affect a natural source of calibration such as  $^{12}\text{B}$  events.  
 2787 The  $^{12}\text{B}$  is a cosmogenically produced isotope through the passage of muons inside the LS. The  $^{12}\text{B}$   
 2788 decays via  $\beta^-$  emissions with a Q value of 13.5 MeV, with more than 98% of the decay resulting  
 2789 in ground state  $^{12}\text{C}$ . This results into the e- energy spectrum shown on Fig. 2.12b. The energy  
 2790 regime involved here is similar to that of reactor IBDs. The  $^{12}\text{B}$  events will be cleanly identified by  
 2791 looking for delayed high-energy  $\beta$  events after an energetic muon. The  $^{12}\text{B}$  events will be uniformly  
 2792 distributed in the detector: moreover,  $^{12}\text{B}$  events are produced continuously, and therefore can be  
 2793 used to follow finally time variations in the detector behavior. As with calibration sources, energy  
 2794 spectra obtained with measured and simulated  $^{12}\text{B}$  events can be compared to control the accuracy  
 2795 of the detector model. This *physics control sample* still presents the disadvantage to be an electron  
 2796 source.

2797 The limitations of the calibration and control samples mentioned above could hide subtle data/MC  
 2798 discrepancies that might be able to bias the results of the oscillation analysis. We fear this problem  
 2799 in particular when ML algorithms are used, due to their ability to use exhaustively the information  
 2800 present in the detector. But, while we have an idea of where the issues could come from, the manual  
 2801 production of event perturbations that go unseen when using these samples would be very time  
 2802 consuming. That's why we propose to use the power of ML for an automated generation of adequate  
 2803 perturbation scenarios. We choose to develop an Adversarial Neural Network (ANN) to produce  
 2804 those perturbations if they exist. A schematic of the concept is presented in Figure 6.4. We try here  
 2805 to extend to a large scale detector the concept introduced in [135].

2806 This network should produce physically plausible perturbations that would not be seen by the  
 2807 calibration system but also by the visualization of the event. If the ANN manages to produce such  
 2808 perturbations, we can derive systematic uncertainties from it. If it fails to find any, it is a proof of  
 2809 robustness for the targeted reconstruction method.

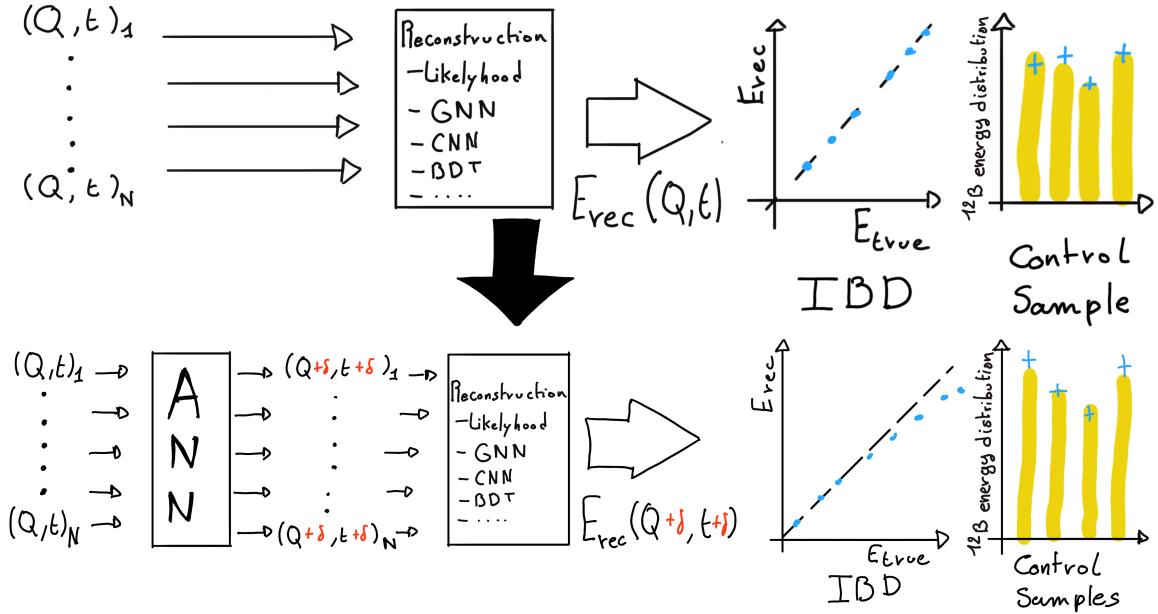


FIGURE 6.4 – Schema of the method to discover vulnerabilities in the reconstruction methods. **On the top** of the image, the standard data flow. The individual charge and times are fed to a reconstruction algorithm. From the reconstructed energies, we can produce an IBD spectrum and compute control observables from the calibration and/or control samples (like a  $^{12}B$  sample). In an ideal case, these observables should include the energy spectrum, the interaction position distributions, and any other useful variables. On the sketch above, the yellow distribution represent a real data sample while the blue points represents a simulated sample. **On the bottom**, the same data flow but we add an ANN between the input and the reconstruction. Still in an ideal case, the ANN learns what slight change to impose to each and every PMT so that the input charge and time so the reconstruction algorithm inaccurately reconstruct the IBD energy, but the perturbation is not visible in the control samples.

### 2810 6.2.1 ANN Architecture

2811 For this study, we consider a “physics” dataset composed of 1M positron events from J23, uniformly  
 2812 distributed in the Central Detector (CD) and in deposited energy between  $E_{dep} \in [1.022; 10.022]$ . This  
 2813 set represents the IBD events we want to the reconstruction to be fooled on.

2814 We use a second “control” dataset of 1M electron events from J23, also uniformly distributed in the  
 2815 detector and over the same energy range. They mimic the energy deposition of  $^{12}B$  decay and are  
 2816 used as the sample to compute the control observables.

2817 This work is a collaboration with an engineer from Subatech Gilles Grasseau. We, the JUNO’s  
 2818 Subatech group, developed the idea and the global method design and the results interpretation.  
 2819 Me and Gilles developed the ANN architecture, and Gilles the FFNN architecture. Gilles was in  
 2820 charge of the Python implementation and I provided the MC samples and readers to read the ROOT  
 2821 data into Python.

2822 We can describe the goal of the ANN by using following loss function:

$$\mathcal{L} = \mathcal{L}_{adv} + \mathcal{L}_{reg} \quad (6.4)$$

2823 where  $\mathcal{L}_{adv}$  is the adversarial loss, which is minimal when the reconstruction is “broken”, i.e. when  
 2824 the delta perturbations introduced on Fig 6.4 are at work. We thus need to define what is a *wrong*

reconstruction. We choose to define it through the correlation between the reconstructed and deposited energy

$$\mathcal{L}_{adv} = |\text{Corr}(E_{rec}, E_{dep})| \quad (6.5)$$

where  $E_{rec}$  and  $E_{dep}$  are the reconstructed energy and the true deposited energy respectively. This loss is positive or null and is minimal when the reconstructed energy after perturbation is decorrelated with the true deposited energy. If this loss is below 1, there is a chance that the delta perturbations imposed on each PMT alter the result of the oscillation analysis. This loss is evaluated on the physics dataset.

The term  $\mathcal{L}_{reg}$  is the regularisation term, which is minimal when the control variables are correctly reconstructed

$$\mathcal{L}_{reg} = \sum_{\lambda} (O_{\lambda}^{rec} - O_{\lambda}^{th})^2 \quad (6.6)$$

where  $\lambda$  index the different control observables that will be considered in this study. It's minimal when the control observables after perturbation  $O_{\lambda}^{rec}$  are coherent with their expected values  $O_{\lambda}^{th}$ . In this exploratory work, we choose as the control observable the difference between the reconstructed position and energy and the ground truth from the Monte Carlo simulation. When this loss is minimal, there is a chance that the perturbations will not be seen in data/MC studies using this control sample, which is one of the goals of the algorithm.

$$\mathcal{L}_{reg} = \sum_{\lambda \in \{x, y, z, E\}} (\lambda_{rec} - \lambda_{true})^2 \quad (6.7)$$

This loss is evaluated on the control dataset.

To these two loss, we adjoin a penalty term  $P$

$$\mathcal{L} = \mathcal{L}_{adv} + \mathcal{L}_{reg} + P \quad (6.8)$$

This penalty  $P$  is here to prevent the ANN from producing event too different from the initial event. It will be further detailed in Section 6.2.4. This loss is evaluated on both datasets.

We see that the final loss is an equilibrium between the adversarial and regularisation loss.

## 6.2.2 Back-propagation problematic

We would like this method to be applicable to any kind of reconstruction algorithm but this is complicated considering standard training method through backward-propagation (method discussed in details in Section 3.1.3) for reasons developed in this section. This force use to develop, in this exploratory work, a new NN for reconstruction. This NN is presented in section 6.2.3.

For explanation, let's define the application of the reconstruction algorithm as  $\mathcal{F}$  on an event  $X$ , resulting in the prediction  $Y$ , and the application of the ANN  $\mathcal{G}$  on  $X$  to give a perturbed event  $X'$ . We can parametrize the equation 6.4

$$Y = \mathcal{F}(X); Y' = \mathcal{F}(X') = \mathcal{F}(\mathcal{G}(X)) \quad (6.9)$$

$$\mathcal{L} \equiv \mathcal{L}(\mathcal{F}(\mathcal{G}(X)), Y_t) \quad (6.10)$$

where  $Y_t$  is the reconstruction target of  $Y$ .

Now if we consider the learnable parameters  $\theta$  of the ANN on which we want to optimize  $\mathcal{L}$ , in the backward-propagation optimisation framework we need to compute

$$\frac{\partial \mathcal{L}(\mathcal{F}(\mathcal{G}(X)))}{\partial \theta} \quad (6.11)$$

2854 which, when using the chain rule, become

$$\frac{\partial \mathcal{L}(\mathcal{F}(\mathcal{G}(X)))}{\partial \theta} = \frac{\partial \mathcal{G}}{\partial \theta} \cdot \frac{\partial \mathcal{F}}{\partial \mathcal{G}} \cdot \frac{\partial \mathcal{L}}{\partial \mathcal{F}} \quad (6.12)$$

2855 The terms  $\frac{\partial \mathcal{G}}{\partial \theta}$  and  $\frac{\partial \mathcal{L}}{\partial \mathcal{F}}$  are easily computable but  $\frac{\partial \mathcal{F}}{\partial \mathcal{G}}$  depends on the nature of the reconstruction  
2856 algorithm.

2857 While this term comes naturally when using neural network algorithms, its computation is embed-  
2858 ded in most of modern framework, it's not so trivial for other types of algorithms like likelihood.  
2859 Solutions exist to optimize networks that work in complex, non-differentiable environments, such  
2860 as *Deep Reinforcement Learning* [136, 137], but as a first prototype, we will restrict ourselves to neural  
2861 networks for the reconstruction algorithm.

2862 The choice to use gradient descent, and therefore neural networks, also allowed us to keep all  
2863 technical software development wrapped in the same language and framework, PyTorch [100].

2864 The backward-propagation introduce a second issue. At the beginning of the subsection we intro-  
2865 duce  $X' = \mathcal{G}(X)$ , the event after perturbation. It's an input of the reconstruction  $\mathcal{F}$ , thus, let's say  
2866 that the event, in its form  $X$ , is a list of tuples  $(id, Q, t)$  which are the hit on the PMT  $id$ . If  $\mathcal{F}$  require  
2867 the information to be formatted in a specific way (graph, images, ...) via an algorithm  $\tau(X)$ , it means  
2868 that

$$\frac{\partial \mathcal{L}(\mathcal{F}(\tau(\mathcal{G}(X))))}{\partial \theta} = \frac{\partial \mathcal{G}}{\partial \theta} \cdot \frac{\partial \tau}{\partial \mathcal{G}} \cdot \frac{\partial \mathcal{F}}{\partial \tau} \cdot \frac{\partial \mathcal{L}}{\partial \mathcal{F}} \quad (6.13)$$

2869 which also requires that  $\frac{\partial \tau}{\partial \mathcal{G}}$  is differentiable.

2870 On the other hand, if  $X$  is already formatted as the input of  $\mathcal{F}$ , it means that  $\mathcal{G}$  takes the same format  
2871 as input, and we drop the requirement on  $\tau$  to be differentiable. Specifically, if  $\mathcal{F}$  takes an image as  
2872 input, it means that  $\mathcal{G}$  will also take an image as input and output an image. Unfortunately, this also  
2873 means that if some information is lost before  $\mathcal{G}$ , for example, during the charge and time aggregation  
2874 in pixels, the ANN cannot retrieve and modify it.

2875 A more elegant solution would that  $\mathcal{G}$  would also compute the transformation  $\tau$  in addition to  
2876 finding relevant perturbation, but for the simplicity of this exploratory work, we use a  $\mathcal{G}$  that process  
2877 transformed data.

### 2878 6.2.3 Reconstruction Network (FFNN)

2879 As introduced just before, we need a NN algorithm for IBD reconstruction. We could have used the  
2880 GNN presented in Chapter 5 but we preferred a more simpler approach to not be constrained by the  
2881 memory consumption of the reconstruction network. The memory issue do not really comes from the  
2882 reconstruction network but from the ANN. The requirement to produce outputs that have the same  
2883 structure and complexity as the reconstruction network makes it even more memory consuming than  
2884 the reconstruction network, thus the choice for a simpler reconstruction network. This network is  
2885 designated as FFNN for "F"-Fully connected Neural Network where "F" is reminder to the  $\mathcal{F}$  from  
2886 previous section.

2887 This network takes as input a vector containing the results of the aggregation of charge and time on  
2888 pixels, forming a vectorized image. We consider JUNO to be composed of 3072 pixels defined by the  
2889 Healpix [127] pixelization. On each of these pixels, we sum the charges and keep the first time of hit,  
2890 resulting in 3072  $(Q, t)$  tuples. To these tuples, we adjoin the position of the center of these pixels,  
2891 resulting in 3072  $(Q, t, x, y, z)$  tuples. The data is finally represented as a  $3072 \times 5 = 15360$  vector. In  
2892 the case where the charge in a pixel is 0, the time is set to 2048 ns, which is way after the closure of  
2893 the trigger window.

2894 The charge is expressed in  $N_{pe}$  and the time of hit in nanoseconds. The time is negative, meaning  
 2895 that 0 ns is the first hit time and -2048 ns is the latest hit time.

2896 FFNN is a Fully Connected Neural Network (FCDNN) composed of the following layers: the input  
 2897 layer, providing the 15360-item vector, followed by fully connected linear layers with the respective  
 2898 number of neurons being [8192, 4096, 2048, 1024, 512, 256, 128, 64, 32]. These layers possess a Leaky  
 2899 ReLU activation function defined as

$$\text{LeakyReLU} = \begin{cases} x, & \text{if } x > 0 \\ 10^{-2} \cdot x, & \text{otherwise} \end{cases} \quad (6.14)$$

2900 The last layer is a linear layer with 4 neurons, representing  $(x, y, z, E)$  without an activation function.

2901 The loss used is the Mean Square Error (MSE)

$$\text{MSE}(\boldsymbol{\eta}, \boldsymbol{\eta}^{true}) = \sum_i (\eta_i - \eta_i^{true})^2 \quad (6.15)$$

2902 where  $\eta$  takes the values of  $(x, y, z, E)$ .

2903 The optimizer used for its training is the Stochastic Gradient Descent with momentum

$$\boldsymbol{\theta}_{t+1} = \boldsymbol{\theta}_t - \Lambda \left( \sum_{i=0} \frac{\partial \mathcal{L}}{\partial \boldsymbol{\theta}_{t-i}} \cdot 0.9^i \right) \quad (6.16)$$

2904 where  $\boldsymbol{\theta}_t$  is vector of learnable parameters at step  $t$ .  $\Lambda$  is the learning rate set at  $10^{-3}$ . The difference  
 2905 with the classical SGD is the gradient term with  $i > 1$ . We save the gradient computed in the previous  
 2906 step and use them as momentum with a decaying weight. The factor 0.9 is an hyperparameter that  
 2907 has been selected for the training.

2908 Additionally, to prevent over-fitting, we introduce a weight decay. Each step, we reduce the amplitude  
 2909 of the parameters  $\boldsymbol{\theta}$  by  $10^{-3}$ :

$$\boldsymbol{\theta}_{t+1} = \boldsymbol{\theta}_t \cdot (1 - 10^{-3}) \quad (6.17)$$

## 2910 Performances

2911 The FFNN is trained independently from the ANN. The dataset is comprised of 1M positrons events  
 2912 uniformly distributed in the detector and in energy over  $E_{dep} \in [1, 10]$  MeV. The training dataset  
 2913 account for 990'000 events with 10'000 events reserved for validation. The data are normalized,  
 2914 mean shifted to 0 and standard deviation scaled to 1, before being processed by the network.

2915 Each epochs goes trough the entire training datasets, with a batch size of 64. The training last for 25  
 2916 epochs. The performance the FFNN are presented in Figures 6.5 and 6.6. We remind that goal of this  
 2917 FFNN is not to have competitive performances against classical algorithms like OMILREC but more  
 2918 to have a simple, NN reconstruction algorithm to run the ANN against.

### 2919 6.2.4 Adversarial Neural Network (ANN)

2920 The ANN aims to introduce perturbations in the event data in such a way that these perturbations  
 2921 are not detectable in the control dataset while still degrading the energy reconstruction of the IBD  
 2922 dataset. For this purpose, and for the reasons detailed in Section 6.2.2, the ANN operates on the  
 2923 inputs of the reconstruction network presented above, namely the FFNN. During the training, the  
 2924 parameters of the FFNN are frozen, meaning they will not be updated during the ANN training. If

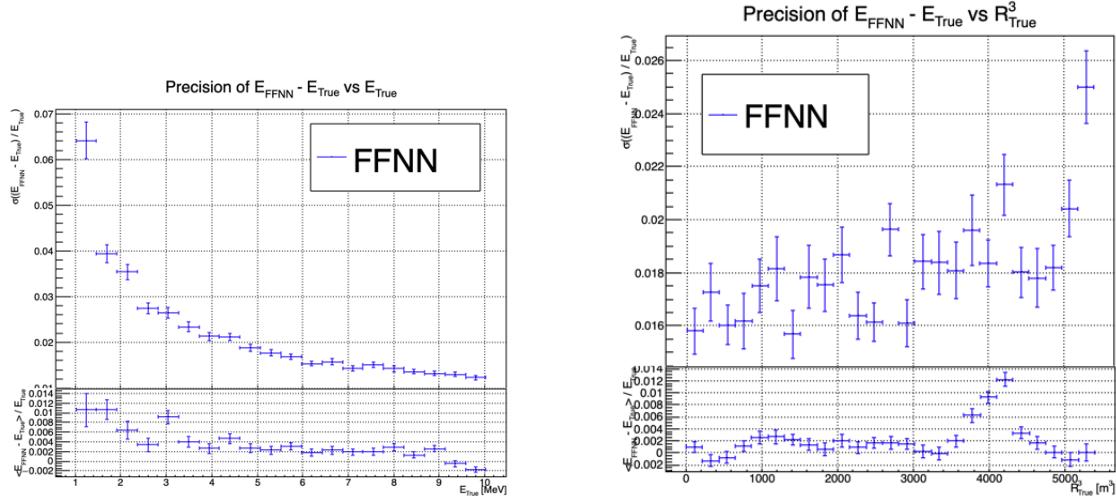


FIGURE 6.5 – Energy resolution of the FFNN with respect to the energy (On the left) and with respect to the radius (On the right)

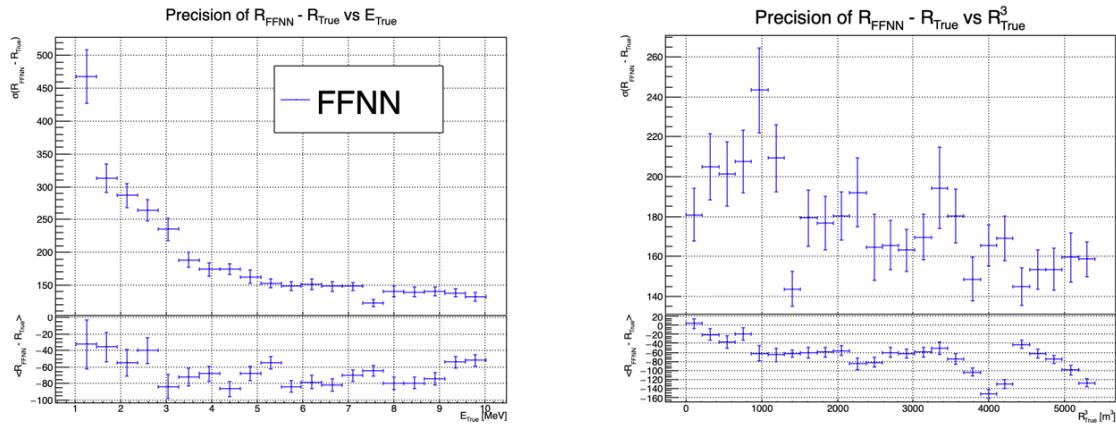


FIGURE 6.6 – Radial resolution of the FFNN with respect to the energy (On the left) and with respect to the radius (On the right)

they were free to be optimized, they would adapt to the perturbations of the ANN, which would go against the objective of this work.

The FFNN takes as input a vector of  $5 \times 3072$  values, representing the  $(x, y, z, Q, t)$  of 3072 Healpix pixels (pixelisation illustrated in Figure 5.2). Those values come from the aggregation of the PMTs belonging to those pixels.

It seems unreasonable that the ANN would modify the Healpix pixel positions, as they are derived from a mathematical construction. It could, however, perturb which PMTs are assigned to specific pixels, introducing localization errors, but the position of the PMTs is carefully monitored during JUNO's construction. Such aggregation errors would likely arise from PMTs located at the edges of the pixels, yet this scenario seems unlikely. Moreover, due to the constraints mentioned in Section 6.2.2, the ANN is required to work with the same format that the FFNN uses as input.

At the start of the project, we attempted to have it operate on both time and charge information simultaneously, but it struggled to converge. After discussions with colleagues in the collaboration, we decided that the ANN would only introduce perturbations in the charge information, as most of

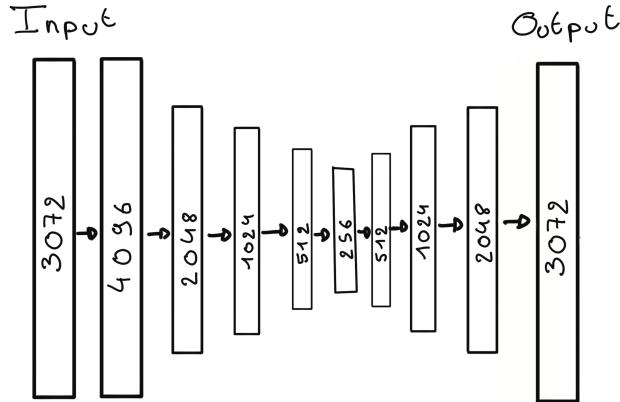


FIGURE 6.7 – Illustration of the “bottleneck” architecture of the ANN. Each block represent a fully connected layer with, on the left, the input layer and on the right the output layer. We see a first reduction of the number of neurons per layer, going from 4096 to 256, followed by an augmentation back to 4096 neurons, thus the “bottleneck”

2939 the energy information comes from the charge.

2940 Our ANN thus needs to output a 3072-dimensional vector, which represents the updated charges of  
2941 the detector.

2942 We decided on a Fully Connected Deep NN (DNN) “bottleneck” architecture for the ANN, illus-  
2943 trated in Figure 6.7. This architecture places a 4096-neuron-wide layer after the input, followed by  
2944 smaller layers of sizes 2048, 1024, and 512 neurons, before finally reaching the 256-neuron layer.  
2945 From this layer, the size increases again to 512, 1024, and finally 2048 neurons before the output  
2946 layer, which consists of 3072 neurons.

2947 The idea behind this architecture is that, by reducing the number of neurons per layer, we force  
2948 the network to summarize the event in 256 parameters, that it will use to regenerate an event. This  
2949 architecture has also the advantage of keeping the number of learnable parameters relatively small,  
2950 as the connection between small layers do not require a lot of parameters.

## 2951 ANN loss

As it was mentioned in the introduction of Section 6.2.1, the loss of the ANN is composed of two losses, the adversarial loss  $\mathcal{L}_{adv}$  and the regularisation loss  $\mathcal{L}_{reg}$ . To those two losses, we adjoin a penalty term that prevent the ANN from producing non-physical events.

$$\mathcal{L} = \mathcal{L}_{adv} + \mathcal{L}_{reg} + P$$

2952 The adversarial loss  $\mathcal{L}_{adv}$  is defined as the absolute value correlation between the reconstructed  
2953 energy and the energy deposit (Eq. 6.5). The regularisation loss  $\mathcal{L}_{reg}$  is the MSE of the true and  
2954 reconstructed energy position vector  $(x, y, z, E)$  (Eq. 6.7).

2955 The penalty term is here to prevent the network from generating event that are too far from the initial  
2956 event. The relevance of this term and its parameters will be further discussed in section 6.2.5. The  
2957 penalty  $P$  is a function that takes the pixelated event  $X$ , its transformation after the ANN  $\mathcal{G}(X)$  and  
2958 a constraint  $\epsilon$

$$P(X, \mathcal{G}(X), \epsilon) = \sum_{i=1}^{3072} (ReLU(-\mathcal{G}(X)_i) + D_i) \quad (6.18)$$

2959 with

$$D_i = \begin{cases} \frac{(X_i - \mathcal{G}(X)_i)^2}{X_i^2} & \text{if } \frac{|X_i - \mathcal{G}(X)_i|}{X_i} > \epsilon \\ 0 & \text{otherwise} \end{cases} \quad (6.19)$$

2960 where  $i$  index the Healpix pixels. The term  $\text{ReLU}(-\mathcal{G}(X)_i)$  is minimal, equal 0, when the charge  
 2961 after perturbation is positive. This term prevent the ANN from producing negative charge, feat  
 2962 impossible for the PMTs.

2963 The second term  $D_i$  is equal to 0 when the relative charge between the original and perturbed pixel  
 2964 is less than  $\epsilon$ . The value of  $\epsilon$  will change during the training, as it will be explained in Section  
 2965 6.2.5. Otherwise, it is the square of this relative charge difference. This term penalize the ANN from  
 2966 producing charges too different from the original event.

2967 When dealing with multiple losses like this, it is important keep them of the same order of magnitude,  
 2968 as we do not want one term to absorb the other.

2969 The loss  $\mathcal{L}_{adv}$  range from 0 to 1 while  $\mathcal{L}_{reg}$  is 0 when the vertex and energy is perfectly reconstructed.  
 2970 It can theoretically go up to infinity. In practice we expect it to take value of the order of magnitude  
 2971 coherent with the reconstruction performances. In fact, if it would take higher value, it would mean  
 2972 that the reconstruction would reconstruct the event far away from the true vertex and energy in  
 2973 comparison to the expected performance. This kind of issue would be immediately be detected,  
 2974 even with simplistic reconstructions such as the charge barycenter, which goes against the goal of  
 2975 producing subtle fluctuation.

2976 We evaluate  $\mathcal{L}_{reg}$  with  $(x, y, z)$  in meter and  $E$  in MeV. If the event is reconstructed with a precision  
 2977 of 15 cm and an energy resolution of 3% at 1 MeV, taking the reconstruction performance of the best  
 2978 reconstruction algorithm OMILREC (see Sections 3.3 and 5.7),  $\mathcal{L}_{reg} \approx 0.3^2 + 0.03^2 = 0.0909$ . We see  
 2979 about an order of magnitude between  $\mathcal{L}_{adv}$  and  $\mathcal{L}_{reg}$ . To compensate for it we weight  $\mathcal{L}_{reg}$

$$\mathcal{L} = \mathcal{L}_{adv} + 60 \cdot \mathcal{L}_{reg} + P(\epsilon) \quad (6.20)$$

2980 The amplitude of  $P$  and the value of  $\epsilon$  will be further discussed in Section 6.2.5.

### 2981 Hyperparameter optimization

2982 All the ANN hyperparameters presented above have been optimized through the numerous iteration  
 2983 the architecture went through. The training is computationally expensive as we need to host both  
 2984 networks on the GPU card, reaching quickly the memory limit of the GPU. The training of the ANN  
 2985 can takes up to 90h. The requirement of having a powerful GPU can be met locally, as Subatech  
 2986 possess an available A100 [122] card with 40GB of memory. We could not port over computing  
 2987 center as they only possess V100 [123] GPU with 20GB of memory.

2988 Those constraint made a random search optimization impossible. It is maybe possible, through  
 2989 optimisation, to reduce the memory requirements to reach the threshold to run on V100 but the  
 2990 challenge was deemed not worth it for an exploratory work.

### 2991 6.2.5 Training of the ANN

2992 The ANN training is divided into two phases. In the first phase, the network learns to accurately  
 2993 reproduce physical events, ensuring that it can handle the intrinsic variability of the detector's re-  
 2994 sponse. This step is crucial, as it provides the foundation for the second phase, where the network  
 2995 searches for subtle perturbations that can degrade the reconstruction without being detected by  
 2996 standard calibration procedures. Splitting the training into two phase also allow to save a version

2997 of the network that know how to reproduce the physical events. We can then “resume” the training  
 2998 from this point if we just introduce changes in Phase 2, saving the training time of Phase 1.

2999 For both phases, we use the both of the datasets presented in section 6.2. We use a batch size of 64  
 3000 for both datasets meaning that, for each steps, the network see 128 events.

3001 Each epochs goes through the entirety of the training dataset.

3002 **First training phase: back to physics**

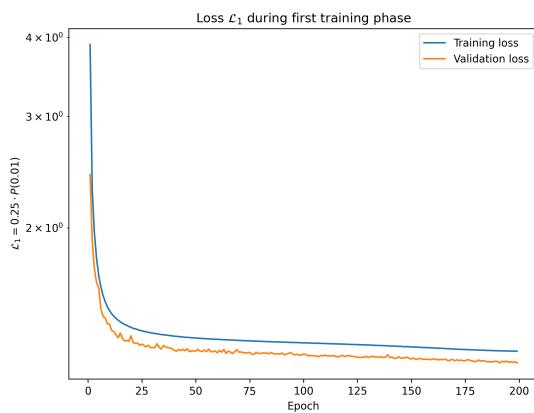


FIGURE 6.8 – Evolution of the loss  $\mathcal{L}_1 = 0.25 \cdot P(0.01)$  during the first phase of the training

3003 When the ANN is initialized, before any training has been done, its parameters are initialized with  
 3004 random values. Multiple initialization methods exist. In this work, we use a common He initializa-  
 3005 tion [138], which is the default initialization in the PyTorch [100] library. If we were to ask for an  
 3006 event from the ANN without training first, the results would be random noise. We thus first have  
 3007 the ANN learn to reproduce physical events.

3008 For this, we conduct a training of 200 epochs where the loss consists only of the penalty term. For  
 3009 scaling purposes, the penalty  $P$  is scaled by 0.25.

$$\mathcal{L}_1 = 0.25 \cdot P(\epsilon = 0.01) \quad (6.21)$$

3010 During this phase, the only objective of the network is to yield events that are the same as the original  
 3011 events.

3012 The evolution of this loss  $\mathcal{L}_1$  during the training for the training dataset and the validation dataset is  
 3013 presented in Figure 6.8. We see that the ANN converges to some stability in the loss.

3014 The time and charge channels of two events, after this training phase, are presented in Figures 6.9  
 3015 and 6.10. We remind that the ANN only act on the charge channel of the event.

3016 We observe that for a localized event, Figure 6.9, the ANN correctly reproduces the event, while  
 3017 for a more diffuse event, Figure 6.10, it produces a more uniform charge distribution. By looking at  
 3018 the color scale in Figure 6.10, we observe that the ANN does not reproduce singular high numbers  
 3019 of  $N_{pe}$ . The highest pixel in the original was 12  $N_{pe}$ , whereas after the ANN, the highest pixel is  
 3020 5  $N_{pe}$ . Furthermore, whereas in the original event the charge repartition, while diffuse, was still  
 3021 concentrated in specific pixels, the ANN spreads the charges in all the pixels.

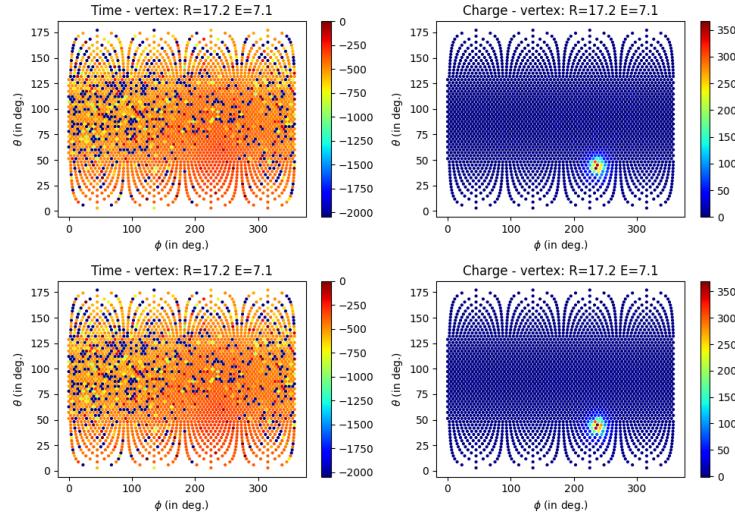


FIGURE 6.9 – Time channel (on the left) and charge channel (on the right) of a **radial, high energy event** ( $R = 17.2$  m,  $E_{dep} = 7.1$  MeV), **Top:** before the ANN perturbation, **Bottom:** after the ANN perturbation. The ANN have been trained for 200 epochs, just after Phase 1. Time channel in ns and charge channel in  $N_{pe}$ .

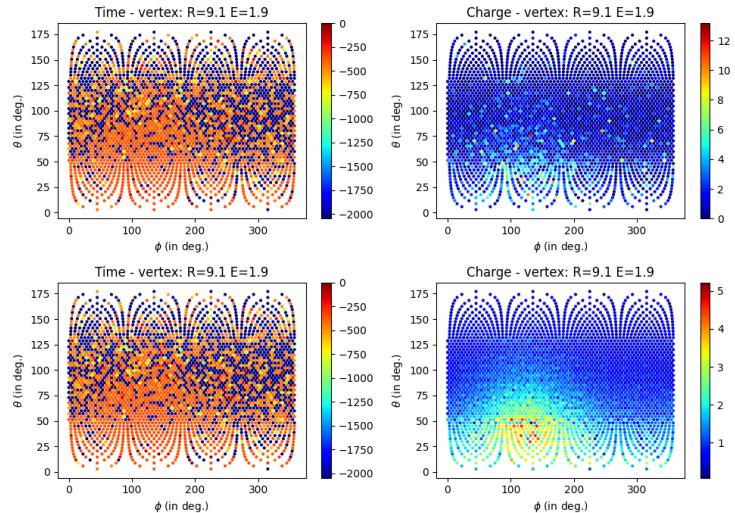


FIGURE 6.10 – Time channel (on the left) and charge channel (on the right) of a **central, low energy event** ( $R = 9.1$  m,  $E_{dep} = 1.9$  MeV), **Top:** before the ANN perturbation, **Bottom:** after the ANN perturbation. The ANN have been trained for 200 epochs, just after Phase 1. Time channel in ns and charge channel in  $N_{pe}$ .

3022 In the next figures, we discuss the reconstruction of the FFNN ( $\mathcal{F}$ ) with and without the presence of  
 3023 the ANN ( $\mathcal{G}$ ) at the end of this Phase 1. The reconstruction by the FFNN of an event perturbed by the  
 3024 ANN is denoted  $(\mathcal{F} \circ \mathcal{G})$ . We differentiate the reconstruction between the two datasets, presented in  
 3025 Section 6.2: the physics dataset, designated as IBD, and the control dataset, designated  $^{12}B$ .

3026 In Figure 6.11, we show the ratio between the reconstructed energy distribution before and after the  
 3027 application of the ANN. For the  $^{12}B$  dataset, the ratio is close to one except in the bin  $E_{rec} > 9.5$  MeV,  
 3028 where we see an excess of events after the ANN. For the IBD dataset, the ratio is close to 1 over the  
 3029 energy range.

3030 In Figure 6.12, we present the distribution of the relative reconstruction errors  $(E_{rec}, E_{dep})/E_{dep}$  with  
 3031 (light histogram) and without (dark histogram) the perturbations predicted by the ANN. We see that  
 3032 without the ANN, the distribution was centered on 0, whereas with it, we observe a small positive  
 3033 bias. In the second row of the histogram, the ratio between the light and dark histograms, we see  
 3034 confirmation of the previous observation, with a deficit of events for  $-0.05 < (E_{rec}, E_{dep})/E_{dep} <$   
 3035 0.02 and an excess of events for  $(E_{rec}, E_{dep})/E_{dep} > 0.02$ . This shift to higher energy explains the  
 3036 excess of events seen in the highest energy bins in Fig. 6.11. The behavior between the  $^{12}\text{B}$  dataset  
 3037 (green histogram) and the IBD dataset (blue histogram) is similar.

3038 At the end of this first reconstruction phase it's apparent that this exploratory ANN is not able to  
 3039 correctly reconstruct the event. This could come from the bottleneck architecture: the reduction of  
 3040 the event to 256 parameters is not enough to correctly rebuild the event. This subject will be further  
 3041 discussed in the conclusion of this chapter in Section 6.3.

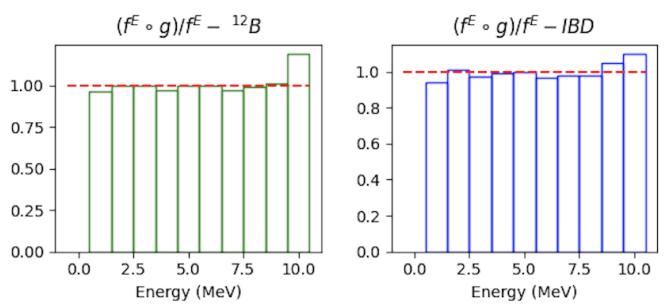


FIGURE 6.11 – Ratio of the reconstructed energy spectra between  $(\mathcal{F} \circ \mathcal{G})$  and  $\mathcal{F}$  at then end of Phase 1 of the training. **On the left :** For the  $^{12}\text{B}$  dataset. **On the right :** For the IBD dataset

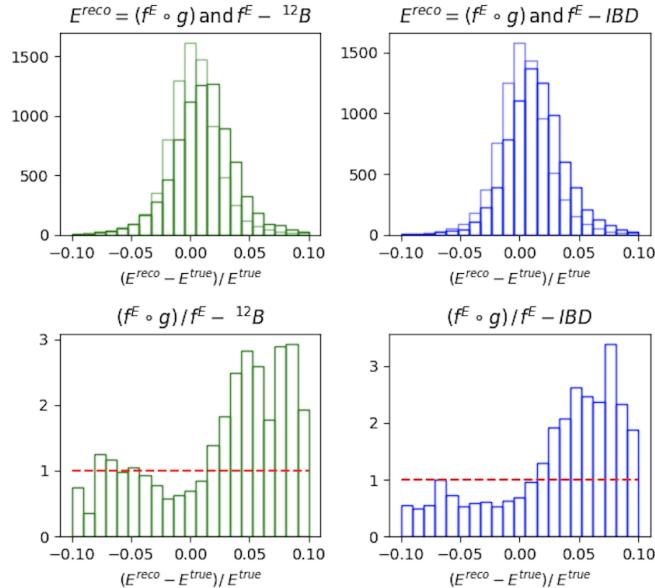


FIGURE 6.12 – **On the top :** Distribution of the relative energy reconstruction error between  $\mathcal{F}$  (light histogram) and  $(\mathcal{F} \circ \mathcal{G})$  (dark histogram) at then end of Phase 1 of the training. **On the bottom :** Ratio between the light and dark histogram of the top figure.

3042 **Second training phase: Breaking of the reconstruction**

3043 Once the ANN is able to reproduce physical events, we change the loss so that it starts to search for  
 3044 potential perturbations. For this we introduce the term  $\mathcal{L}_{adv}$  and  $\mathcal{L}_{reg}$  producing a second loss  $\mathcal{L}_2$ .  
 3045 Adding those terms will significantly change the loss. The previous minima in the parameter phase  
 3046 space the ANN found minimizing  $\mathcal{L}_1$  will not be the minima  $\mathcal{L}_2$ . To prevent a gradient explosion,  
 3047 we introduce a growing factor  $\lambda$  in front of the term  $\mathcal{L}_{adv}$  and  $\mathcal{L}_{reg}$ . This factor starts at  $\lambda = 0.01$  at  
 3048 epoch 201 and grows  $\lambda_{i+1} = \lambda_i + 0.01$  where  $i$  indexes the epoch. It caps at  $\lambda_{max} = 1$  at epoch 300  
 3049 after which it stops growing.

3050 Also to ease the task of the ANN, we relax the constraint in the penalty term  $P$  from  $P(0.01)$  to  
 3051  $P(0.15)$ .

3052 The expression of the phase 2 loss  $\mathcal{L}_2$  becomes:

$$\mathcal{L}_2 = \lambda (\mathcal{L}_{adv} + 60 \cdot \mathcal{L}_{reg}) + 0.25 \cdot P(0.15) \quad (6.22)$$

3053 This second phase of the training last for 200 more epochs, up to epoch 400.

3054 The profiles of  $\mathcal{L}_2$ ,  $\mathcal{L}_{adv}$ ,  $60 \cdot \mathcal{L}_{reg}$  and  $0.25 \cdot P(0.15)$  during this second phase of the training are  
 3055 presented in Figures 6.13 and 6.14. The profile of the loss  $\mathcal{L}$  over entirety of the training is presented  
 3056 in figure 6.15.

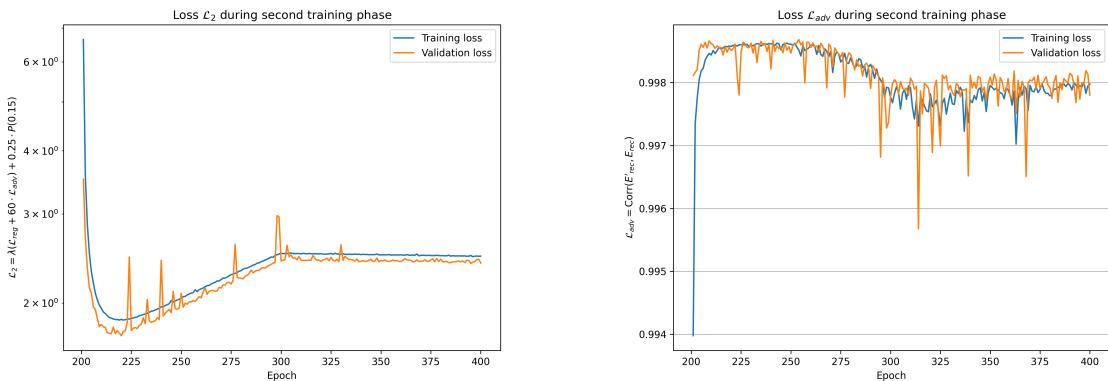


FIGURE 6.13 – Profile of the loss  $\mathcal{L}_2$  and  $\mathcal{L}_{adv}$  during the second phase of training. The linear increase of  $\mathcal{L}_2$  is due to the growing factor  $\lambda$  in Eq. 6.22.

3057 We see on Figures 6.13 and 6.14 that during the first epochs of this second phase,  $L_{reg}$  and  $P(0.15)$   
 3058 decrease fast. At the end of the 1st phase, the work to recover the initial reconstruction seems  
 3059 incomplete since the performance of FFNN were not recovered (Fig. 6.12). During the first epochs  
 3060 of the second phase,  $L_{reg}$  seems to continue this work. This is logical since this term is suppose to  
 3061 temperate the perturbations so they are not visible with a control sample. At the same time, we see  
 3062 a quick increase of  $L_{adv}$ , confirming that the reconstruction is less broken.

3063 During most of the next 50 epochs of this phase, all losses are more stable before  $L_{adv}$  starts decreasing,  
 3064 between epochs 250 and 300. It suggests  $L_{adv}$  has managed to re-deteriorate the reconstruction,  
 3065 despite the quasi stability (or slow decrease) of  $L_{reg}$ . This is the desired behavior concerning the  
 3066 losses.

3067 Unfortunately, as can be seen in Figs. 6.18 and 6.12, the performance of the reconstruction is still too  
 3068 deteriorated : it would very probably be detected by data/MC comparisons with control samples.  
 3069 Also, on these figures and on Figure 6.20, we can't see an indication that IBD events are affected  
 3070 more than  $^{12}B$  events by the perturbation. A difference here is not mandatory : the same deterioration

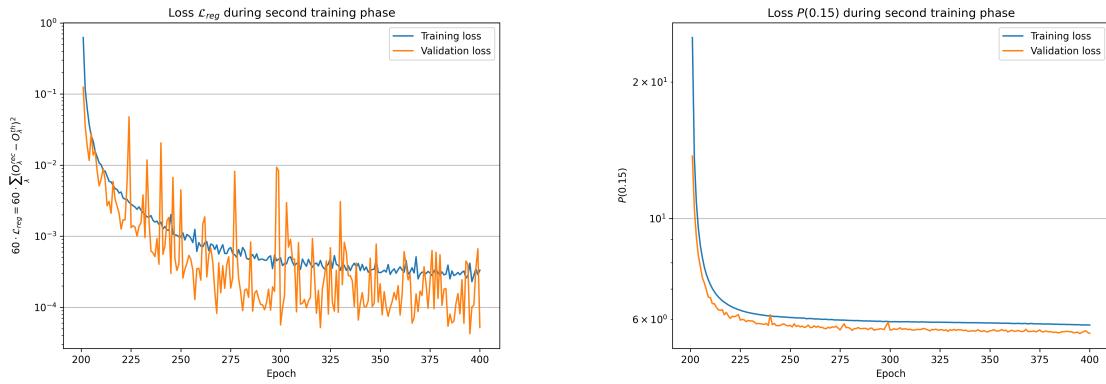


FIGURE 6.14 – Profile of the loss  $60 \cdot L_{reg}$  and  $0.25 \cdot P(0.15)$  during the second phase of training

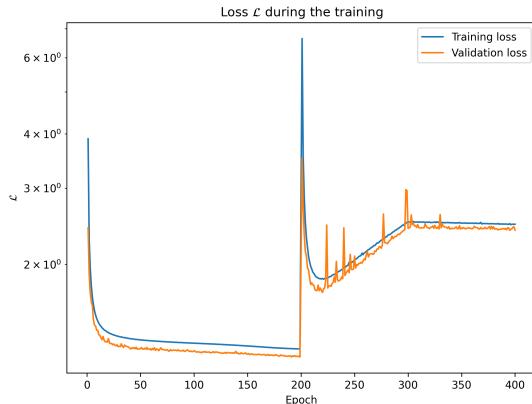


FIGURE 6.15 – Profile of the loss over the entirety of the training (Phase 1 and 2)

3071 of the resolution could still be undetectable in data/MC comparison of the distributions of the energy  
 3072 or of the vertex position in  $^{12}B$  samples, and still be enough to alter the oscillation analysis. However,  
 3073 observing a difference here would have been a good sign.

3074 After 200 epochs of Phase 2, the correlation in  $\mathcal{L}_{adv}$  is still at 0.998, the penalty term  $P(0.01)$  is stable  
 3075 and the regularisation loss  $\mathcal{L}_{reg}$  is close to stability.

3076 For illustration, events produced by the ANN after 400 epochs are displayed in Figures 6.16 and 6.17.  
 3077 These are the same event as displayed in Figures 6.9 and 6.10.

3078 The same observations that were made after phase 1 still apply after phase 2. The ANN still spreads  
 3079 the charge over multiple pixels for central events, Figure 6.17, while for radial events it is able to  
 3080 reproduce the small localization of the event.

3081 When looking at the distribution of ratio between the reconstructed energy distribution before and  
 3082 after the application of the ANN, Figure 6.18, we observe this time a deficit of events in the high  
 3083 energy bin. This deficit is explained by the comparison between the distribution of relative recon-  
 3084 struction errors, Figure 6.19, in which we see a small negative bias. This same figure shows a wider  
 3085 loss in resolution when the ANN is present. This is the ANN working to degrade the resolution of  
 3086 the FFNN.

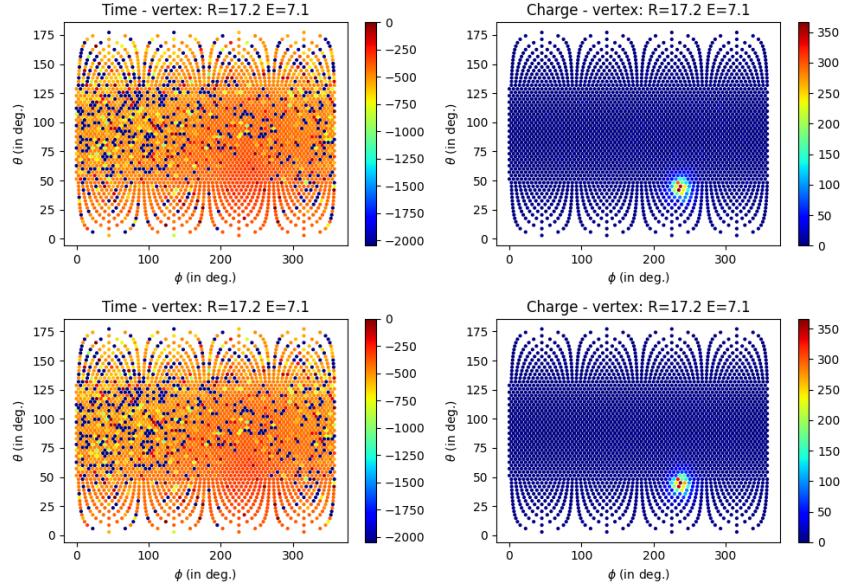


FIGURE 6.16 – Time channel (on the left) and charge channel (on the right) of a **radial, high energy event** ( $R = 17.2$  m,  $E_{dep} = 7.1$  MeV), **Top:** before the ANN perturbation, **Bottom:** after the ANN perturbation. The ANN have been trained for 400 epochs, just after Phase 2. Time channel in ns and charge channel in  $N_{pe}$ .

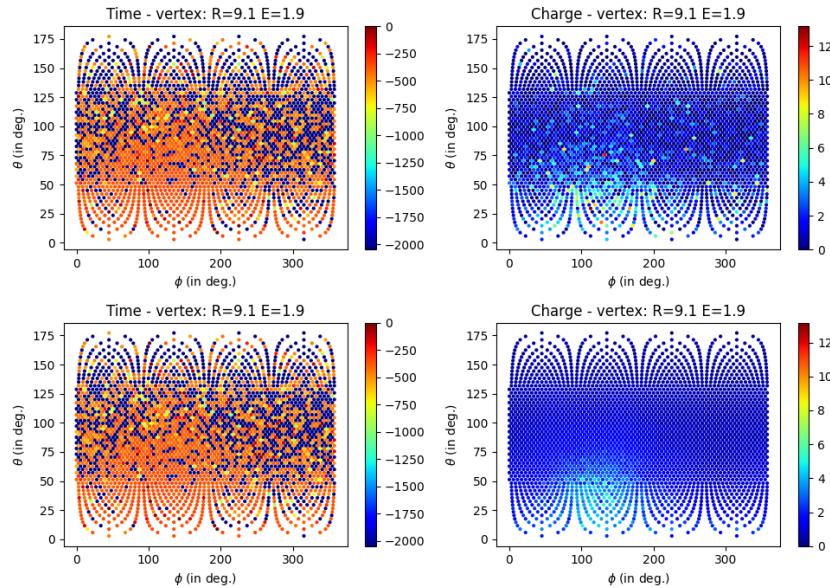


FIGURE 6.17 – Time channel (on the left) and charge channel (on the right) of a **central, low energy event** ( $R = 9.1$  m,  $E_{dep} = 1.9$  MeV), **Top:** before the ANN perturbation, **Bottom:** after the ANN perturbation. The ANN have been trained for 400 epochs, just after Phase 2. Time channel in ns and charge channel in  $N_{pe}$ .

3087 Figure 6.20 shows the ratio between the relative error on the reconstructed energy between the IBD  
 3088 and the  $^{12}\text{B}$  dataset with and without the ANN. We don't see any indicative difference, the ANN  
 3089 even seems to have harmonized the reconstruction error between the two datasets.

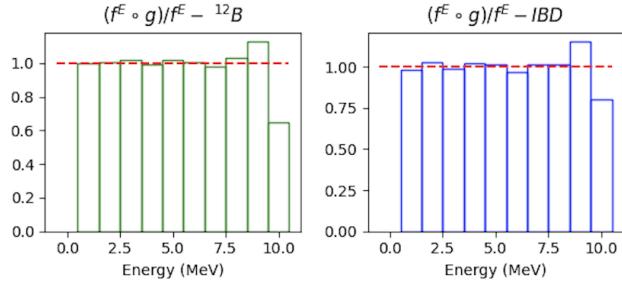


FIGURE 6.18 – Ratio of the reconstructed energy spectra between  $(\mathcal{F} \circ \mathcal{G})$  and  $\mathcal{F}$  at the end of Phase 2 of the training. **On the left :** For the  ${}^{12}B$  dataset. **On the right :** For the IBD dataset

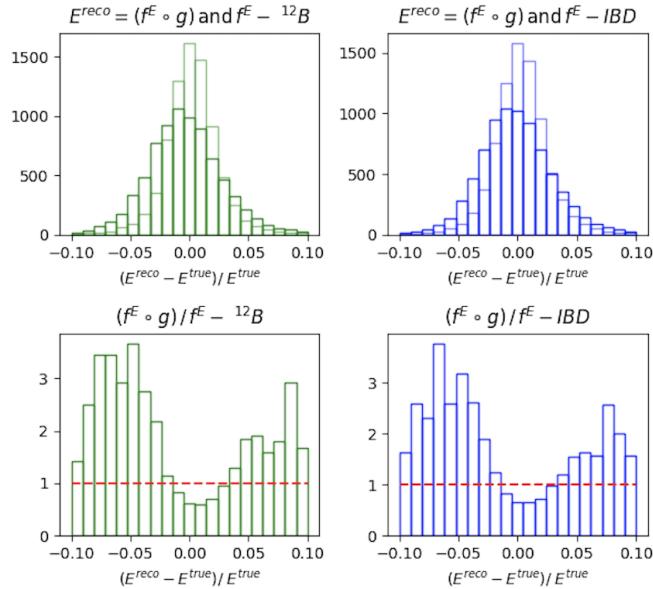


FIGURE 6.19 – **On the top :** Distribution of the relative energy reconstruction error between  $\mathcal{F}$  (light histogram) and  $(\mathcal{F} \circ \mathcal{G})$  (dark histogram) at the end of Phase 2 of the training. **On the bottom :** Ratio between the light and dark histogram of the top figure.

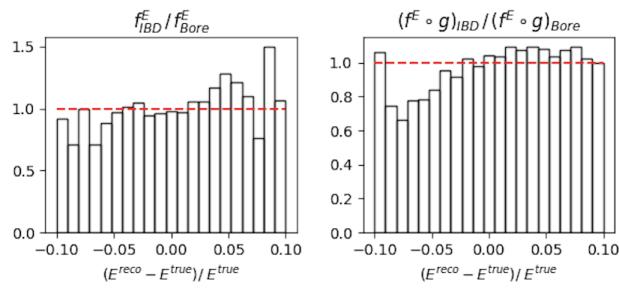


FIGURE 6.20 – Ratio between the relative error on the reconstructed energy between the IBD and the  ${}^{12}B$  dataset. **On the right :** without the ANN. **On the left :** with the ANN.

3090 In the next section, we will summarize the lessons we gathered while working on this ANN, as well  
 3091 as some perspectives for the future.

### 3092 6.3 Conclusion and prospects

3093 Reliability and knowledge of our reconstruction algorithms are crucial for the successful conduct  
 3094 of the experiment. The first step to testing and comparing the reconstruction algorithms is to have  
 3095 them publicly available. To this end, I have implemented a BDT for energy reconstruction in JUNO's  
 3096 common software and compared its performance and behavior in detail to the classic likelihood algo-  
 3097 rithm OMILREC. The strong correlation between their errors indicates that close to no improvement  
 3098 can be made by combining the two algorithms, as they use the same information.

3099

3100 Concerning the development of an ANN, we actually developed prototype which was useful to  
 3101 identify several of the difficulties we will have to overcome in the future to produce an ANN fulfilling  
 3102 the aims defined at the beginning of this chapter. First, we determined that learning individual  
 3103 perturbations for each of the 17600 LPMTs (meaning more than 35000 learnable parameters if one  
 3104 decides to perturb  $Q$  and  $t$ ) is too much for our available hardware. Then, we determined that par-  
 3105 ticular techniques would be necessary to solve the back propagation problems described in Section  
 3106 [6.2.2](#) and couple the ANN with any reconstruction algorithm pre-existing in JUNO.

3107 After having opted for a simpler prototype with 3072 learnable parameters, we faced the problem  
 3108 due to their random initialization and adapted by adding a new term to the loss function and splitting  
 3109 the training in 2 phases. Then, we experience the importance of the definition of the loss function,  
 3110 in particular the way to balance two antagonist terms : one ( $\mathcal{L}_{adv}$ ) which deteriorates the event  
 3111 reconstruction and one that preserves it ( $\mathcal{L}_{reg}$ ). The way to define each of these two terms is not  
 3112 trivial either. Having these notions in mind is essential before trying to develop a tool producing  
 3113 realistic perturbation patterns at the individual PMT level.

3114 With this prototype, we manage to produce one of the desired behaviors for such an ANN : we  
 3115 observe  $\mathcal{L}_{adv}$  deteriorating the reconstruction and  $\mathcal{L}_{reg}$  reducing the deterioration. However, at the  
 3116 end of the training, the deterioration is too high compared to the subtle scenarios we want the ANN  
 3117 to produce. A solution against this could be to give a higher weight to  $\mathcal{L}_{reg}$  or to find a more efficient  
 3118 penalty term  $P(\epsilon)$ .

3119 A smarter definition of  $\mathcal{L}_{adv}$  could also be useful: a definition more explicitly related to our goal  
 3120 (biasing the oscillation analysis) might induce smaller perturbations. This could be a  $\mathcal{L}_{adv}$  favoring  
 3121 a small energy dependent bias between  $E_{rec}$  and  $E_{dep}$  in IBD events. A smarter  $\mathcal{L}_{adv}$  could also help  
 3122 to produce perturbations that affect the IBD reconstruction more than the reconstruction of  $^{12}B$  or  
 3123 calibration events. Although this feature is not mandatory, it would be welcome. It is not achieved  
 3124 by the present version of the ANN.

3125 A possible explanation is the perturbations it produces seem to follow a random pattern across the  
 3126 3072 pixels. It could be the result of the limited efficiency of the first phase of the training (which tries  
 3127 to recover from the random initialisation of these perturbations), or result from the present definition  
 3128 of  $\mathcal{L}_{adv}$  (which can be minimized by random perturbations).

3129 The architecture of the ANN is, for now, very simple; it's a Fully Connected Deep NN with a  
 3130 bottleneck architecture. Previous work in developing ML for reconstruction [114] and the algorithms  
 3131 presented in Chapters 4 and 5 show the relevance of convolutions in the reconstruction, and the work  
 3132 of Gavrikov et al. [115] presented at the beginning of this chapter hints at the importance of the time  
 3133 and charge distribution. A more complex and refined architecture can probably be more effective.

3134 Another architecture improvement could come from ResNet architectures [96]. They have already  
3135 proven that the introduction of residual operations helps the network reach better performance. We  
3136 can imagine a network where instead of  $X' = \mathcal{G}(X)$  we have  $X' = \mathcal{G}(X) + X$ , where the ANN  $\mathcal{G}$   
3137 computes only the perturbation instead of a whole new event.

3138 To eventually design an ANN able to perturb several variables for each and every PMT instead of  
3139 3072 pixels, we need to find a way to reduce the number of learnable parameters. For instance, the  $\delta Q$   
3140 perturbation could be a function common to all PMT, depending on learnable parameters controlling  
3141 its variation against  $Q$ ,  $t$  and the position of the PMT. The choice of the function could also be guided  
3142 by physics informed considerations. To help to learn perturbations affecting IBD events more than  
3143 others, the function could also depend on the initial reconstruction of the interaction position, since  
3144 the reconstruction of annihilation gammas depends on it.

3145 Finally, to use this method on every reconstruction algorithm, we must move away from the back-  
3146 propagation method, for reasons detailed in Section 6.2.2, and use different methods such as Rein-  
3147 forcement Learning.

<sup>3148</sup> **Chapter 7**

<sup>3149</sup> **Dualcalorimetric analysis with  
neutrino oscillation for Precision  
Measurement**

<sup>3151</sup>

<sup>3152</sup> “We demand rigidly defined areas of doubt and uncertainty!”

*Douglas Adams, The Hitchhiker’s Guide to the Galaxy*

<sup>3153</sup>

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<sup>3178</sup>

<sup>3179</sup> JUNO is a high-precision neutrino oscillation experiment. To resolve the Neutrino Mass Ordering (NMO) with the required statistical significance, JUNO must be sensitive to the subtle spectral phase shift, on the order of a few percents, as illustrated in Figure 7.1. This phase shift manifests as a small

<sup>3180</sup><sup>3181</sup>

difference between the Normal Ordering (NO) and Inverted Ordering (IO) spectra, which becomes even smaller after accounting for detection effects such as energy resolution smearing, non-linear detector responses, and background contamination, as shown in Figure 7.2.

This chapter is based on simulated data due to the unavailability of real JUNO data, which will only be available in 2025. The purpose of this analysis is to validate the methods and tools developed for dual calorimetry and neutrino oscillation measurements, ensuring that they are robust and ready for future real data.

Among other condition, a precise and complete understanding of the reconstruction and detector effects is crucial. The challenge reside in the technology used in the detector, which, while based on well known technology: scintillator observed by PMT, is being deployed on a scale never seen before, in term of scintillator volume and PMT size. Understanding every effects that goes in the detector can become extremely complicated. Any method to help detecting problems is therefore welcome. Comparing the data and results obtained by two systems measuring the same events, but subject to different sources of error, is therefore precious. This is the purpose of the dual calorimetry techniques used in JUNO thanks to the existence of 2 PMT systems: the LPMT and SPMT systems.

The reconstruction of the IBD positron energy must be very performant: an unprecedented resolution of 3% at 1 MeV [53] is necessary to determine the NMO with the aimed significance.

Furthermore, an energy scale uncertainty below 1% is essential to accurately assess the likelihood of the NO and IO hypotheses. If this uncertainty exceeds 1%, systematic biases could distort the reconstructed spectra, potentially leading to the erroneous exclusion of the correct mass ordering hypothesis (NO or IO). For instance, a shift in the energy scale could mimic a phase shift between the spectra, making it possible to wrongly favor NO when IO is true, or vice versa. This effect has been studied in the introduction of Chapter 4 of [82].

Understanding all the effects influencing the detector response can be quite complex. Consequently, any methodologies that facilitate problem detection and validation of the reconstruction processes are essential for ensuring accurate results.

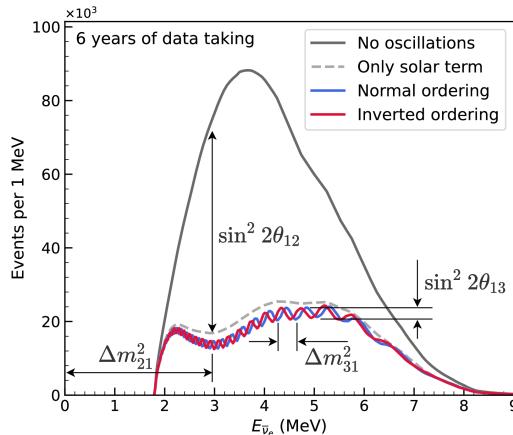


FIGURE 7.1 – Expected number of neutrinos event per MeV in JUNO after 6 years of data taking. The black curve shows the flux if there was no oscillation. The light gray curve shows the oscillation if only the solar terms are taken in account ( $\theta_{12}$ ,  $\Delta m_{21}^2$ ). The blue and red curve shows the spectrum in the case of, respectively, NO and IO. The dependency of the oscillation to the different parameters are schematized by the double sided arrows. We can see the NMO sensitivity by looking at the fine phase shift between the red and the blue curve.

One detector effect to take into account is the detector non linearity. Detector non-linearity can introduce significant biases in the energy reconstruction of events, compromising the precision of

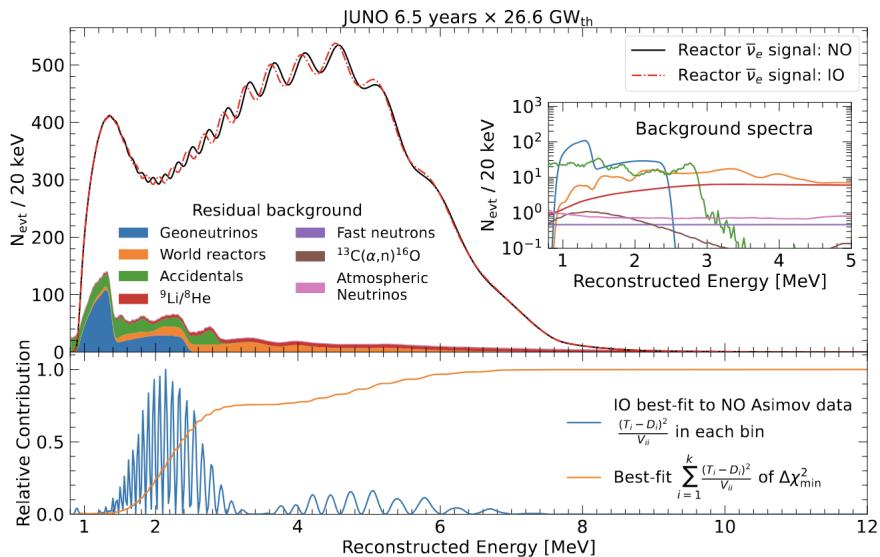


FIGURE 7.2 – Oscillated reactor  $\bar{\nu}_e$  spectra for the Normal Ordering (Black) and Inverted Ordering (Red) for 6.5 years data taking and a resolution of 3% without any statistical or systematic fluctuation. Figure from [90].

neutrino oscillation measurements and increasing systematic uncertainties, which could potentially distort the determination of the neutrino mass.

One of the possible source of non-linearity, which will be used as a reference in this chapter, is the charge non-linearity (QNL) that will be discussed in next section. Several dual calorimetry techniques can address this issue. Some are calibration techniques, that are also described in section 4.3 of [82]. More generally, comparing the results of the two systems will allow for the detection of potential issues on the calibration or reconstruction. This is done in this thesis by comparing directly the spectra and oscillation parameters measurements of the two PMT systems. We call this kind of dual calorimetry "Dual calorimetry with neutrino oscillation", since it is based on the visible energy spectra used by the oscillation analysis of reactor antineutrinos.

In this chapter, we explore several ways to perform this comparison. One of them relies on the difference between the values of  $\Delta m_{21}^2$ ,  $\sin^2(2\theta_{12})$  measured with the LPMT and the SPMT systems. Both systems measure them with similar uncertainties. For reasonable values of the QNL, we expect these differences to be smaller than the individual uncertainties. However, the significance of these differences might still be high. Indeed, both systems reconstruct the same events, therefore the same distribution of the true positron energy, as well as the same scintillation photon emission. Therefore, the energy spectra reconstructed by the two systems share a part of their fluctuations. This translates into correlated reconstructed spectra and consequently lead to correlations between the measurements of  $\Delta m_{21}^2$  and  $\sin^2(2\theta_{12})$ . The uncertainty on the SPMT-LPMT difference is largely decreased by this correlation. Other ways to perform the comparison (see next sections) all rely the reconstructed spectra, therefore on the evaluation of the correlation between the LPMT and SPMT spectra.

In the next section we will discuss the motivations behind this study. In Section 7.2, I present the methods we propose to implement Dual calorimetry with neutrino oscillation, and of the way we estimate their sensitivity. In section 7.3, I present the fit framework used, and then, in section 7.4 the technical improvement brought and the difficulties faced during the development. To end this chapter I present the results in 7.7 and discuss the conclusions and perspectives in 7.8.

## 3240 7.1 Motivations

### 3241 7.1.1 Discrepancies between the SPMT and LPMT results

3242 As mentioned earlier, the SPMT and LPMT systems are expected to detect the same events. There-  
 3243 fore, after proper calibration, any significant discrepancies between the two systems' results could  
 3244 indicate a calibration error, a systematic effect, or an unaccounted detector issue. Detecting such  
 3245 differences is critical, as even small deviations from the expected response could compromise the  
 3246 determination of the Neutrino Mass Ordering (MO) or introduce systematic biases in the oscillation  
 3247 parameter measurements, leading to incorrect conclusions about the true mass ordering.

3248 Both systems are anticipated to show similar sensitivity to the oscillation parameters  $\theta_{12}$  and  $\Delta m^2_{21}$   
 3249 [61]. Therefore, any detected discrepancies will be based on these parameter measurements. A sim-  
 3250 ple comparison of the values and independent uncertainties from the two systems could highlight  
 3251 discrepancies. However, we believe – and will demonstrate in this chapter – that an independent  
 3252 analysis of each system lacks critical information. By considering both statistical and systematic  
 3253 correlations between the two systems, we can design more robust and powerful statistical tests.

3254 Our work in this chapter is to develop such tools, which in practice implies to define test statistics. A  
 3255 first step will be to determine the distribution of these test statistics in the case when no unexpected  
 3256 problem affects the LPMT nor the SPMT problem. This will give us the distribution of those statistical  
 3257 test in absence of discrepancies. Later, the value of the test statistics that we will measure in real data  
 3258 can be compared to these distributions to produce p-values, to judge of the potential present of an  
 3259 unexpected effect.

3260 To evaluate the power of our methods, we need to simulate a concrete difference between the two  
 3261 spectra. We have chosen to study a specific potential effect, Charge Non-Linearity (QNL), which will  
 3262 be detailed in the following section. QNL affects the reconstructed energy spectrum by introducing  
 3263 a non-linear relationship between the true and measured charge in the PMTs. Our statistical tests  
 3264 are designed to detect such distortions, and they should be sensitive to unexpected effects –such as  
 3265 calibration errors or insufficient simulation precision – as long as the induced distortion exceeds a  
 3266 threshold of approximately 1-2% in the reconstructed energy spectrum.

### 3267 7.1.2 Charge Non-Linearity (QNL)

3268 The energy response of the Central Detector (CD) is influenced by two types of non-linearity. The  
 3269 first arises from the intrinsic properties of the Liquid Scintillator (LS), where the photon production  
 3270 is not linearly proportional to the deposited energy, as shown in Figure 2.12a. This non-linearity  
 3271 results from a combination of scintillation and Cherenkov light production. The scintillation yield is  
 3272 governed by Birk's law, which introduces a "quenching" effect that depends on the particle type and  
 3273 energy. Additionally, Cherenkov radiation, which constitutes less than 10% of the collected light,  
 3274 introduces a velocity-dependent non-linearity. These physical non-linearities in the LS contribute to  
 3275 the overall non-linearity of the energy response before any further distortions from the photomulti-  
 3276 plier tubes (PMTs)

3277 The second type of non-linearity comes from the LPMT charge measurements. When photons hit a  
 3278 PMT and give rise to PEs, a current pulse is formed. In the photon counting regime, simply exceeding  
 3279 a certain threshold allows to conclude that a single photon hit the PMT. When several photons hit the  
 3280 PMT simultaneously, one enters the photon integration regime : the pulse is sampled and integrated  
 3281 over a certain time window to produce a reconstructed charge Q. Calibration methods are applied  
 3282 to determine the relationship between the charge Q and the number of PEs (which is the quantity  
 3283 proportional to the energy deposit one wants to measure). Several effects impact this procedure:  
 3284 the signal pulse can fluctuate and be distorted between two events where the same number PEs  
 3285 occurred; the PMT gain might not be linear as a function of the number of photons that hit the PMT;

3286 the charge reconstruction algorithm is not supposed to be perfect, and its results are further affected  
 3287 by electronic noise and inter-channel cross-talk. The impact of these effects grows with the number  
 3288 of PEs.

3289 Precedent studies, Section 4.2.3 of [82], suggest a model for the channel-wise QNL:

$$\frac{Q_{rec}}{Q_{true}} = \frac{-\gamma_{qnl}}{9} Q_{true} + \frac{\gamma_{qnl} + 9}{9} \quad (7.1)$$

3290 where  $Q_{rec}$  is the reconstructed number of PE by the PMT,  $Q_{true}$  is true number of PE that hit the  
 3291 PMT, and  $\gamma_{qnl}$  is a factor representing the amplitude of the non-linearity.

3292 Studies at previous experiments, like Daya Bay, concluded that the best reachable control of QNL  
 3293 in the 1-10 PEs range was  $\gamma_{qnl} = 0.01$  [139]. As already mentionned in Section 2.3.2, JUNO LPMTs  
 3294 operate in a larger range : 1-100 PEs (See also table 7.1). In such a case, a realistic value of  $\gamma_{qnl}$  is not  
 3295 known.

	1PE	2~5PE	5~10PE	10~20PE	20~50PE	50~100PE	>100PE
LPMT	42.56%	40.54%	8.74%	5.12%	2.80%	0.24%	0.003%
SPMT	95.19%	4.80%	0.01%	0%	0%	0%	0%

TABLE 7.1 – The charge fraction in terms of the number of PE collected at the single  
 PMT for the reactor  $\bar{\nu}_e$  IBD events. Table taken from [82]

3296 The event-wise impact resulting from the channel-wise QNL can be parameterised this way :

$$\frac{E_{vis}^{rec}}{E_{vis}^{true}} = \frac{-\alpha_{qnl}}{9} E_{vis}^{true} + \frac{\alpha_{qnl} + 9}{9} \quad (7.2)$$

3297 In JUNO, the visible energy is proportional to the number of emitted photons per unit energy deposit.  
 3298 It includes the physical non linearities. In the equation above  $E_{vis}^{true}$  is this visible energy, while  $E_{vis}^{rec}$   
 3299 is what it becomes when the reconstructed charges found in an event are modified according to Eq.  
 3300 7.1.

3301 An example is shown on Fig. 2.14, where we show the  $E_{vis}^{rec}/E_{vis}^{true}$  ratio for several samples of  
 3302 uniformly distributed electron events, generated with various values of  $E_{vis}^{true}$ . Here, an extreme  
 3303 value  $\gamma_{qnl} = 0.05$  was assumed. On can see on Fig. 2.14 that it corresponds to a 2% effect at 8 MeV,  
 3304 equivalent to  $\alpha_{qnl} = 0.025$ . The effect of Eq 7.2 is illustrated in Figure 7.3.

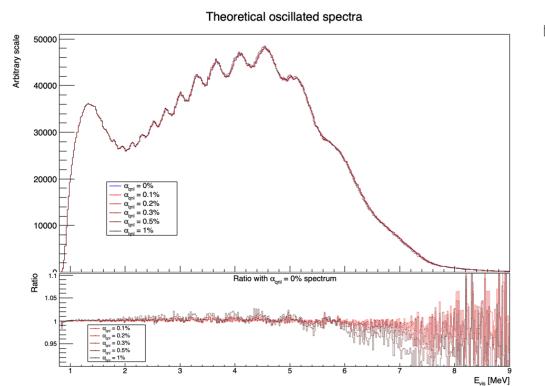


FIGURE 7.3 – On top: Oscillated spectra for different value of  $\alpha_{qnl}$ . On bottom: Ratio  
 of the number of event with  $\alpha_{qnl} = 0\%$ .

3305 This example is from references [82], which aimed at demonstrating the potential of the dual

3306 calorimetry calibration method mentioned in section 2.4.3. If it works as hoped, the residual event-  
 3307 wise QNL effect will be below 0.3%. In this chapter, we propose methods to detect residuals higher  
 3308 than this.

3309 Fig. 7.5 show several other examples with varying  $\gamma_{qnl}$  values, and the corresponding values of  $\alpha_{qnl}$ .  
 3310 Using 1M events from the JUNO official simulation J23.0.1-rc8.dc1 (released on 7th January 2024), we  
 3311 simulated events up to the photon collection in LPMTs and introduced an additional channel-wise  
 3312 QNL by using the equation 7.1 to modify the number of collected photons.

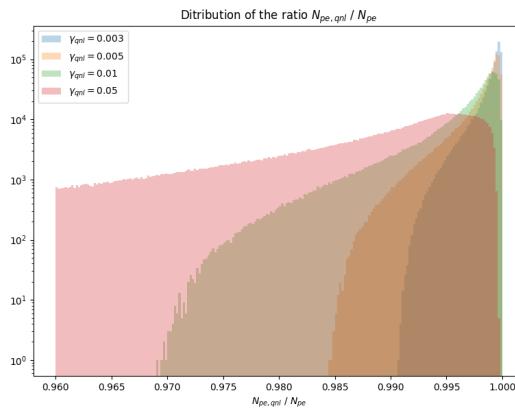


FIGURE 7.4 – Distribution the ratio reconstructed charge (in nPE equivalent) over the number of collected nPE for different value of  $\gamma_{qnl}$ . We use a sample of 1 million positron event uniformly distributed in the detector and in energy in the range  $E_{dep} \in [1, 10] \text{ MeV}$

3313 In Figure 7.4 we show the distribution of the ratio of the reconstructed charge (in nPE equivalent)  
 3314 over the number of collected nPE for different values of  $\gamma_{qnl}$ . The right parts of those distribution,  
 3315 where the ratio is close to 1, are mostly central events. The charge is homogeneously distributed, the  
 3316 effect of the channel-wise QNL is reduced because the PMTs each collect a relatively small number  
 3317 of nPE. The left tail, with ratio < 1, are radial events, the photons are concentrated in a small number  
 3318 of PMTs, the effect of the channel wise QNL is greater.

3319 In Figure 7.5, we show the mean of the distributions of Figure 7.4 as a function of the energy. From  
 3320 the 8.5 MeV data point, we compute an effective  $\alpha_{qnl}$ . The effect of this effective  $\alpha_{qnl}$  is represented  
 3321 as the dashed line. On the bottom of Fig 7.5 is presented the charge ratio difference between the  
 3322 effective  $\alpha_{qnl}$  and the mean effect of a  $\gamma_{qnl}$ . We see that the event-wise QNL, described by Eq. 7.2,  
 3323 do not represent correctly the channel-wise QNL described by Eq. 7.1 at low energy. Indeed, Eq. 7.2  
 3324 assume no QNL effect at 1 MeV, where in reality some of the PMTs will still suffer from QNL.

3325 Despite this difference, the necessity to use the effective event-wise model expressed by Eq. 7.2,  
 3326 and consequently to find the correspondence between values of  $\gamma_{qnl}$  and  $\alpha_{qnl}$ , instead of directly the  
 3327 channel wise model of Eq. 7.1 will be explained in Section 7.2.1.

## 3328 7.2 Our approach to Dual Calorimetry with neutrino oscillation

3329 In this section, we describe 4 statistical tests that we propose to use to detect unexpected effects in  
 3330 one of the PMT systems. Each test is based on a particular test statistics. In practice, the main result  
 3331 we want to produce in this chapter is the distributions followed by these test statistics.

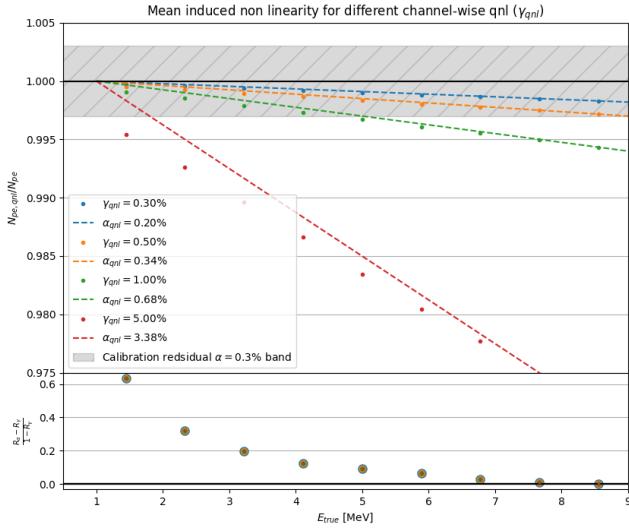


FIGURE 7.5 – **On top:** Ratio of the reconstructed charge (in nPE equivalent) over the number of collected nPE. The dots represent the mean of the distributions in Figure 7.4 and the dashed line are the equivalent event-wise non-linearity from eq 7.2. The hatched zone is the residual non-linearity expected after calibration [84]. **On bottom:** Difference between QNL induced by an event wise QNL and the mean QNL induced by a channel wise QNL. The value for  $\alpha_{qnl}$  and  $\gamma_{qnl}$  follow the color code of the top figure. For a given energy, all the data point have the same value.

In this section, we propose four distinct statistical tests designed to detect unexpected discrepancies between the LPMT and SPMT systems. Each test aims to evaluate different aspects of the reconstructed energy spectra:

1. Test 1 compares the measurements of solar oscillation parameters  $\sin^2 2\theta_{12}$  and  $\Delta m_{21}^2$  derived independently from each system.
2. Test 2 directly compares the LPMT and SPMT spectra bin by bin.
3. Test 3 involves a joint fit of the two spectra, with and without a hypothesis of discrepancy.
4. Test 4 examines the residuals in the fit of oscillation parameters  $\sin^2 2\theta_{12}$  and  $\Delta m_{21}^2$  after the joint fit. The primary objective of this analysis is to establish the distributions of these test statistics under both the null hypothesis (no unexpected effect) and the alternative hypothesis (presence of a discrepancy).

The distributions of these test statistics cannot be analytically determined and are instead generated empirically through toy experiments. In each toy experiment, we generate two spectra of the IBD visible energy: one from the LPMT system and the other from the SPMT system. Since both systems observe the same events, their statistical fluctuations are correlated. To account for this, we compute a  $(820 \times 820)$  covariance matrix that captures both the bin-to-bin correlations within each spectrum and the cross-correlations between the LPMT and SPMT spectra. Details of the sample generation process are provided in Section 7.3.3. Note that we use toy samples rather than samples produced by the full simulation of JUNO since the latter option would not be affordable in terms of computing time.

In the next subsection, we present the informations the reader must know about these spectra to understand the test statistics presented in the rest of the current section.

### 3354 7.2.1 Toy experiments

3355 The sensitivity of our tests depends on the sample size, which scales with the duration of exposure  
 3356 to the antineutrino flux: 100 days, 1 year, 2 years, and 6 years. For each exposure time, we generate  
 3357 1000 toy experiments, where the number of events in the LPMT and SPMT spectra is drawn from a  
 3358 Poisson distribution with the expected mean value for that exposure. Since the same physical events  
 3359 are reconstructed by both systems, their fluctuations are not independent, and we account for the  
 3360 statistical correlations between the LPMT and SPMT spectra in our toy generation process. It was  
 3361 recently evaluated in the recent reference paper on JUNO's sensitivity [90] that about 95000 IBDs  
 3362 would be selected in 6 years.

3363 An example of pair of spectra is shown on Figure 2.16, in the form of two joint histogram of 410, 20  
 3364 keV wide bins each. This is the format used in the fit performed by the present version of the reactor  
 3365 oscillation analysis developed at Subatech. It is important to notice that the IBD events present in  
 3366 the LPMT spectrum of a toy experiment are the same as those in the SPMT spectrum: the same  
 3367 events are just reconstructed twice, by either system. The LPMT and SPMT spectra are therefore not  
 3368 independent : Their respective fluctuations in the number of entries per bin are correlated. These  
 3369 correlations stem from what is common between the LPMT et SPMT spectra, namely :

- 3370 — The statistical fluctuations of the true  $E_{vis}$  distribution (before any reconstruction).
- 3371 — The fluctuation of the number of photons produced by scintillation or Cherenkov effect.

3372  
 3373 When generating toy experiment, the fluctuations drawn in each bin around the average expected  
 3374 number of events must account for these correlations. We therefore evaluated the  $(820 \times 820)$   
 3375 covariance matrix describing the uncertainty on the number of entries in each of the 410 bins of  
 3376 the 2 spectra, as well as the bin-to-bin correlations, especially those between the bins of the LPMT  
 3377 spectrum and those of the SPMT spectrum. This is described in Section 7.5. Here, we just want  
 3378 to emphasize the importance of this point, one of the original tasks to be carried out for the work  
 3379 presented in this chapter.

3380 As already stated earlier, toy experiments will be used to evaluate the distributions of the four test  
 3381 statistics. We will first produce reference distributions: the ones that rule the possible values of  
 3382 the test statistics if none of the PMT systems is affected by any unexpected effect. These references  
 3383 are sufficient to run a test once JUNO will take data: the values of the test statistics obtained in  
 3384 a real data sample can be compared with the reference distributions, to evaluate to which extent  
 3385 the null hypothesis (no unexpected effect) is credible (p-values, or any pertinent quantities, can be  
 3386 computed). This is true whatever the nature of the unexpected effect.

3387 To give an idea of the power of the method, an explicit scenario must be simulated for the un-  
 3388 expected effect. For that purpose, we also generate sets of toy experiments where the  $E_{vis}$  spectrum  
 3389 reconstructed by the LPMT is distorted using Eq. 7.2. We will test the following levels of QNL:  $\alpha_{qnl} \in$   
 3390  $\{0.003, 0.002, 0.001\}$ . As a reminder, the calibration guarantees a residual event-wise non-linearity of  
 3391  $\alpha_{qnl} \leq 0.003$  [84].

3392 The most probable values in the distributions of the test statistics obtained in such cases will be  
 3393 compared with the reference distributions to derive a "median" predicted p-value. One can also  
 3394 compute the probability to observe in real data a p-value lower than a certain value, if the assumed  
 3395 QNL effect actually exists in these data.

3396 When we initiated this work, the best test statistics to use was not obvious to us. This is why we  
 3397 decided to test 4 test statistics, of growing complexity. We present them in the 4 next subsections.

### 3398 7.2.2 Comparing the solar parameters from individual analyses : LPMT vs SPMT

3399 The first test statistics is probably the most natural one: it's essentially a direct comparison of the  
 3400 values of  $\sin^2(2\theta_{12})$  and  $\Delta m_{21}^2$  measured by separate analyses of the LPMT and the SPMT spectra.  
 3401 These analyses are performed using the oscillation fit tool developed at Subatech, described in  
 3402 Sections 2.7 and 7.3. A fit to the LPMT spectrum provides  $\sin^2(2\theta_{12})_L$  and  $\Delta m_{21,L}^2$ , while a separate  
 3403 fit to the SPMT spectrum provides  $\sin^2(2\theta_{12})_S$  and  $\Delta m_{21,S}^2$ .

The direct comparison proceeds in practice via the differences between the fit results :

$$\Delta\theta = \sin^2(2\theta_{12})_L - \sin^2(2\theta_{12})_S \quad (7.3)$$

$$\Delta D = \Delta m_{21,L}^2 - \Delta m_{21,S}^2 \quad (7.4)$$

3404

3405 A very simple test statistics would be for instance

$$S = \frac{|\Delta\theta|}{\sigma_{\Delta\theta}} \quad (7.5)$$

3406 directly related to the significance of the difference between the SPMT and LPMT results. This  
 3407 requires to determine the uncertainty  $\sigma_{\Delta\theta}$ . This cannot be considered as the mere quadratic sum  
 3408 of the uncertainties on  $\sin^2(2\theta_{12})_L$  and  $\sin^2(2\theta_{12})_S$  returned by the fitter. Indeed, because of the  
 3409 correlations, described in the previous subsection, between the LPMT and SPMT spectra, the fitted  
 3410 parameters are also correlated.

3411 The calculation of  $\sigma_{\Delta\theta}$  must account for it. Simple error propagation dictates :

$$\sigma_{\Delta\theta}^2 = \sigma_{\sin^2(2\theta_{12})_L}^2 + \sigma_{\sin^2(2\theta_{12})_S}^2 - 2\sigma_{\sin^2(2\theta_{12})_L}\sigma_{\sin^2(2\theta_{12})_S}C_{L,S} \quad (7.6)$$

3412 where  $C_{L,S}$  is the correlation between the SPMT and LPMT measurements. We expect it to be high  
 3413 (well above 0.9, see Figures 7.6, 7.7, 7.8 and 7.9). Consequently, we expect it to considerably lower  
 3414 the value of  $\sigma_{\Delta\theta}^2$ , and increase the significance  $S$ .

3415 This simple example can be seen as an illustration of the fact that the correlations between the LPMT  
 3416 and SPMT spectra boosts the sensitivity of our test statistics to unexpected effects. Indeed, with 6  
 3417 years of data, and counting only the statistical uncertainties, we expect the statistical uncertainties  
 3418  $\sigma_{\sin^2(2\theta_{12})_L}^2$  and  $\sigma_{\sin^2(2\theta_{12})_S}^2$  to both be around 0.15% [61]. A preliminary evaluation [117] of the impact  
 3419 of an uncorrected QNL effect with  $\alpha_{qnl} = 1\%$  on the value of  $\sin^2(\theta_{12})$  predicted a bias of 0.1%,  
 3420 therefore of 0.05% on  $\sin^2(2\theta_{12})$ . With no correlation, this would lead to a significance  $S$  far below 1.  
 3421 Accounting for the correlation allows far better.

3422 The test statistics we actually use for this direct comparison is a generalisation of the simple one  
 3423 above : it includes both the results on  $\sin^2(2\theta_{12})$  and  $\Delta m_{21}^2$  :

$$\chi_{ind}^2 = \Delta_{ind}^T U^{-1} \Delta_{ind} \quad (7.7)$$

3424 where  $\Delta_{ind}$  is a vector defined as

$$\Delta_{ind} = [\Delta\theta, \Delta D] \quad (7.8)$$

3425 using equations 7.3 and 7.4.

3426 The covariance matrix  $U$  is a  $(2 \times 2)$  matrix containing the uncertainties on the components of  $\Delta_{ind}$   
 3427 and the correlation between them. We derive this matrix from the  $(4 \times 4)$  covariance matrix  $V$ ,  
 3428 which contains the uncertainties on the fitted values of  $\sin^2(2\theta_{12})_L$ ,  $\sin^2(2\theta_{12})_S$ ,  $\Delta m_{21,L}^2$  and  $\Delta m_{21,S}^2$ ,  
 3429 as well as the correlations between these quantities. For that purpose, we simply use the linear error

propagation formalism, that can be found in section 40.2.6 of the statistical review of the PDG 2020 [92] :

$$U = A V A^T \quad (7.9)$$

where the transfer matrix  $A$  is obtained this way

$$A_{ij} = \frac{\partial \Delta_i^{ind}}{\partial \lambda_j} \quad (7.10)$$

where  $\lambda_j$  one of the parameters ( $\Delta m_{21,L}^2, \sin^2(2\theta_{12})_L, \Delta m_{21,S}^2, \sin^2(2\theta_{12})_S$ ). Assuming this indexing order for  $j$  and  $i$  ordering following Eq 7.8,  $A$  is expressed

$$A = \begin{bmatrix} 1 & 0 & -1 & 0 \\ 0 & 1 & 0 & -1 \end{bmatrix} \quad (7.11)$$

We acknowledge that linear error propagation is valid when all fluctuations or uncertainties are gaussian. However, since our results will be based on distributions of  $\chi^2_{ind}$  produced with toy samples, this choice remains valid.

An important ingredient here is to determine the correlation coefficients in  $V$ . On a dedicated set of 1000 toy experiments, we perform fits to the LPMT and SPMT spectra, and compute the correlations empirically from the 1000 sets of best fit values of the solar parameters :  $\sin^2(2\theta_{12})_L$  vs.  $\sin^2(2\theta_{12})_S$ ,  $\Delta m_{21,L}^2$  vs  $\Delta m_{21,S}^2$ ,  $\sin^2(2\theta_{12})_L$  vs.  $\Delta m_{21,S}^2$ , etc. We need the correlations corresponding to the null hypothesis and therefore use toy experiments produced with no QNL effect.

The correlations between these parameters for 100 days, 1 year, 2 years and 6 years can be found in Figures 7.6, 7.7, 7.8 and 7.9 respectively.

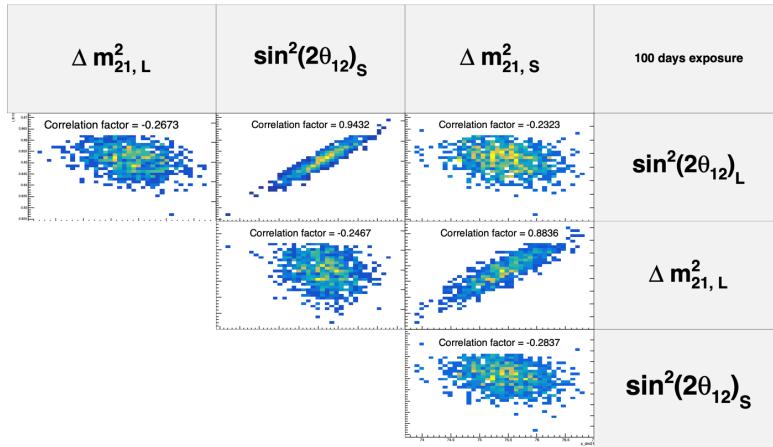


FIGURE 7.6 – Distribution and correlation between the best fit point of 1000 individual toys fit for 100 days exposure without supplementary QNL.

We observe strong correlation between the reconstructed  $\Delta m_{21}^2$  and  $\sin^2(2\theta_{12})$  of both systems as presented in Table 7.2, row one and two. As the relative statistical uncertainty decrease with exposure, the correlations grow ranging from 0.88 to 0.95 for  $\Delta m_{21}^2$  and from 0.94 to 0.98 for  $\sin^2(2\theta_{12})$ . We observe between parameters of the same fit, a small anti-correlation of about -0.25, line 4 and 5 of Table 7.2.

Because the parameters are heavily correlated between the LPMT and SPMT fit, and that  $\Delta m_{21}^2$  and  $\sin^2(2\theta_{12})$  are slightly anti-correlated in the same fit, the couples of different parameters from different fit,  $\text{Corr}(\sin^2(2\theta_{12})_L, \Delta m_{21,S}^2)$  and  $\text{Corr}(\sin^2(2\theta_{12})_S, \Delta m_{21,L}^2)$ , are also anti-correlated.

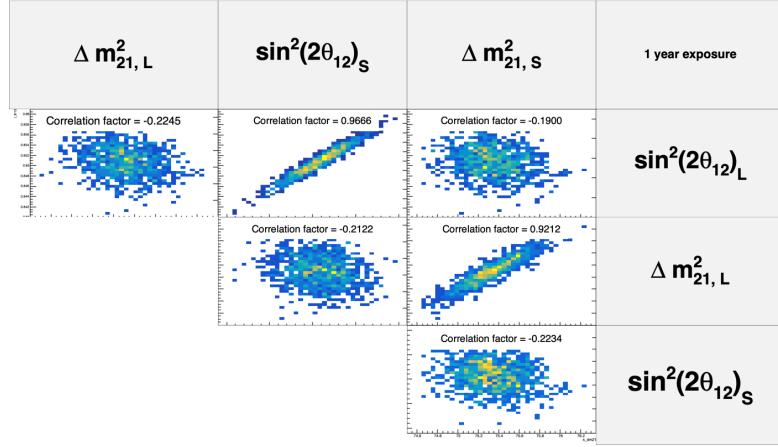


FIGURE 7.7 – Distribution and correlation between the best fit point of 1000 individual toys fit for 1 year exposure without supplementary QNL.

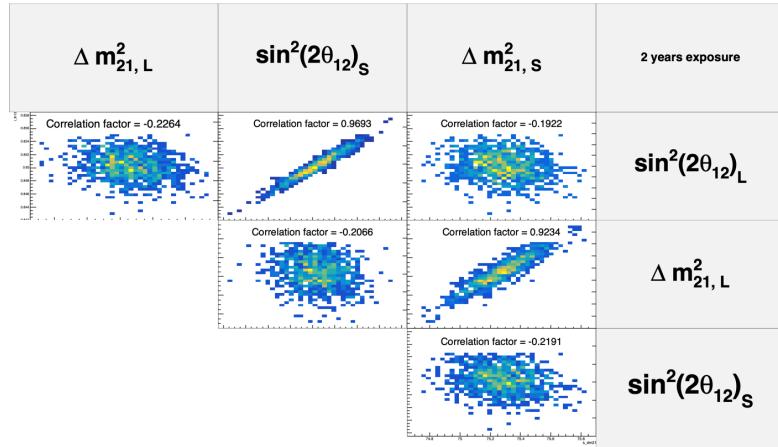


FIGURE 7.8 – Distribution and correlation between the best fit point of 1000 individual toys fit for 2 years exposure without supplementary QNL.

3453 The distributions  $\chi^2_{ind}$  will be Shown in Section 7.7.

### 3454 7.2.3 Direct comparison between the SPMT and LPMT spectra

3455 In the second test, we perform a bin-by-bin comparison of the LPMT and SPMT spectra without  
 3456 fitting any oscillation parameters. Again, we use here a  $\chi^2$ -like statistics. We do not expect the  
 3457 reference distribution (for  $\alpha_{qnl} = 0$ ) to be centered around the number of degree of freedom (i.e. the  
 3458 number of bins of each spectrum in our case) as should be distributed (if the spectra contain enough  
 3459 events in each bin to assume a gaussian behavior of the number of entries) the  $\chi^2$  comparing 2 his-  
 3460 tograms when they are consistent with each other. Indeed, even in the absence of unexpected events,  
 3461 the LPMT and SPMT are quite different because of the very different reconstruction resolutions. We  
 3462 therefore need here again to establish this reference distributions with toys. And compare them later  
 3463 with the distributions obtained for the various tested values of  $\alpha_{qnl}$ .

3464 Our test statistics is :

$$\chi^2_{spe} = \Delta_{spe}^T U^{-1} \Delta_{spe} \quad (7.12)$$

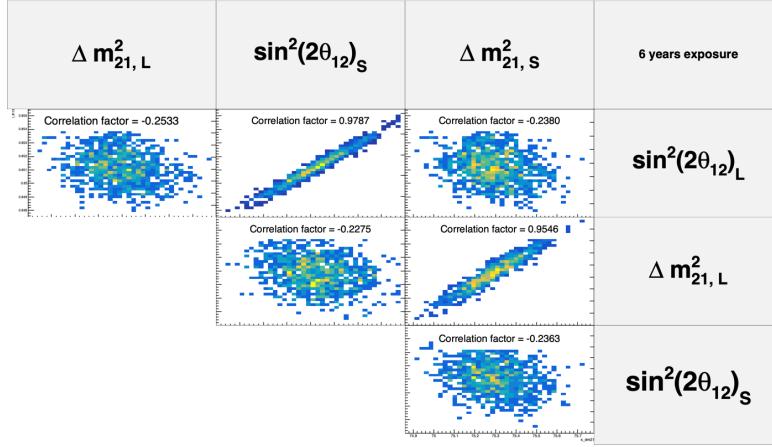


FIGURE 7.9 – Distribution and correlation between the best fit point of 1000 individual toys fit for 6 years exposure without supplementary QNL.

	100 days	1 year	2 years	6 years
Corr( $\Delta m_{21,L}^2, \Delta m_{21,S}^2$ )	0.8836	0.9212	0.9234	0.9546
Corr( $\sin^2(2\theta_{12})_L, \sin^2(2\theta_{12})_S$ )	0.9432	0.9666	0.9693	0.9787
Corr( $\sin^2(2\theta_{12})_L, \Delta m_{21,L}^2$ )	-0.2673	-0.2245	-0.2264	-0.2533
Corr( $\sin^2(2\theta_{12})_S, \Delta m_{21,S}^2$ )	-0.2837	-0.2234	-0.2191	-0.2363
Corr( $\sin^2(2\theta_{12})_L, \Delta m_{21,S}^2$ )	-0.2323	-0.19	-0.1922	-0.2380
Corr( $\sin^2(2\theta_{12})_S, \Delta m_{21,L}^2$ )	-0.2467	-0.2122	-0.2066	-0.2275

TABLE 7.2 – Correlations between the parameters BFP of the individual LPMT and SPMT fits for multiple exposures using 1000 toys.

3465 where

$$\Delta_i^{spe} = h_{L,i} - h_{S,i} \quad (7.13)$$

3466 and

$$U = AVA^T \quad (7.14)$$

3467 Here,  $i$  runs over the 410 bins of the individual spectra. Also,  $h_{L,i}$  and  $h_{S,i}$  are the contents of the  $i$ th  
3468 bin of the LPMT and SPMT spectra respectively. We need to know the uncertainty on  $\Delta_i^{spe}$  and the  
3469 correlations with  $\Delta_j^{spe}$ 's in other bins. We derive them from  $V$ , the  $(820 \times 820)$  covariance matrix  
3470 introduced at the beginning of this section, which can be seen as the covariance matrix of a 820-bin  
3471 double spectrum juxtaposing the LPMT and SPMT spectra. We remind its determination will be  
3472 presented in Section 7.5. To obtain  $U$  from  $V$ , we again apply the linear error propagation, with the  
3473 transfer matrix :

$$A_{ij} = \frac{\partial \Delta_i^{spe}}{\partial h_j} = \frac{\partial (h_{L,i} - h_{S,i})}{\partial h_j} \quad (7.15)$$

3474 Thus,  $A_{ij} = 1$  if  $i = j$ , and  $A_{ij} = -1$  if  $j$  is the SPMT bin corresponding to the  $i$  LPMT bin.

3475 We expect this statistics to have a certain power since  $\chi^2_{spe}$  can be increased for 2 reasons in case of  
3476 unexpected problem: first, the LPMT spectrum (if the LPMT is affected) will be distorted and become  
3477 less consistent with the SPMT spectrum; second, the correlations between the LPMT and SPMT might  
3478 also modified. Since  $V$  present a peculiar correlation pattern (see Section 7.5), a departure from this  
3479 pattern also has some valuable impact on  $\chi^2_{spe}$ .

### 3480 7.2.4 Joint fit of the SPMT and LPMT spectra : $\chi^2_{H_0} - \chi^2_{H_1}$

3481 This kind of fit has already been introduced in Section 2.7. As a reminder, it involves the minimisa-  
 3482 tion of

$$\chi^2_{joint} = (\mathbf{T}(\boldsymbol{\theta}, \mathbf{h}) - \mathbf{D})^T V^{-1} (\mathbf{T}(\boldsymbol{\theta}, \mathbf{h}) - \mathbf{D}) + \ln(|V|) \quad (7.16)$$

3483 where  $\mathbf{T}(\boldsymbol{\theta}, \mathbf{h})$  is the predicted joint LPMT+SPMT spectrum and  $\mathbf{D}$  the corresponding data vector.  
 3484 The matrix  $V$  is the full  $(820 \times 820)$  covariance matrix which incorporate both the statistical uncer-  
 3485 tainties and the bin-to-bin correlations between the LPMT and SPMT spectra.

3486 In this fit, we include the usual oscillation parameters,  $\sin^2(2\theta_{12})$ ,  $\Delta m_{21}^2$ ,  $\sin^2(2\theta_{13})$  and  $\Delta m_{31}^2$  along  
 3487 with two additional parameters,  $\delta \sin^2(2\theta_{12})$  and  $\delta \Delta m_{21}^2$  which allow for a potential discrepancy in  
 3488 the LPMT reconstruction or calibration.

3489 Several remarks must be made here to better understand what we do precisely.

- 3490 — Given JUNO's lack of sensitivity to  $\sin^2(2\theta_{13})$ , this parameter is fixed in the fit to the PDG value  
 3491 (see table 7.3). In most of JUNO's fit procedures (see Section 2.7), it's allowed to float during  
 3492 the minimisation, but is treated like a nuisance parameter, by adding a penalty term based on  
 3493 the PDG central value and uncertainty.
- 3494 — The oscillation fit that we perform here does not really aim at the oscillation parameters in  
 3495 themselves, but is performed to detect a difference between the LPMT and SPMT spectra.  
 3496 JUNO is supposed to be very sensitive to  $\Delta m_{31}^2$  via the LPMT spectrum. However, it has been  
 3497 shown by studies carried out at Subatech (and confirmed since then by other groups in the  
 3498 Collaboration), that up to 2 years of data taking, the presence of multiple minima in  $\Delta m_{31}^2 \chi^2$   
 3499 profile can make its determination delicate. Since  $\Delta m_{31}^2$  is not the aim of our present study,  
 3500 we stabilize the fit by treating this parameter as a nuisance parameter, adding to  $\chi^2_{joint}$  the  
 3501 following penalty term :

$$\chi^2_{\Delta m_{31}^2} = \frac{(\Delta m_{31}^2 - \overline{\Delta m_{31}^2})^2}{\sigma_{\overline{\Delta m_{31}^2}}^2} \quad (7.17)$$

3502 We define two hypothesis. The hypothesis  $H_0$  assumes that no unexpected effect is present, meaning  
 3503 that  $\delta \sin^2(2\theta_{12}) = 0$  and  $\delta \Delta m_{21}^2 = 0$ , and the hypothesis  $H_1$  where  $\delta \sin^2(2\theta_{12}) \neq 0$  and  $\delta \Delta m_{21}^2 \neq 0$   
 3504 are needed to account for any potential calibration or reconstruction bias. The test statistic is then  
 3505 defined as the difference between the minimized  $\chi^2$  values under  $H_0$  and  $H_1$ :

$$\Delta \chi^2 = \chi^2_{joint, H0} - \chi^2_{joint, H1} \quad (7.18)$$

3507 where  $\chi^2_{joint, H0}$  is the result of the minimisation when the fit assumed no unexpected effect (fixing  
 3508  $\delta \sin^2(2\theta_{12})$  and  $\delta \Delta m_{21}^2$  to 0), while  $\chi^2_{joint, H1}$  assumes a possible effect, letting this parameters free  
 3509 to float. A large value of  $\Delta \chi^2$  would indicate a significant deviation from the null hypothesis (no  
 3510 discrepancy), suggesting the presence of an unexpected effect in the LPMT system.

3511 Distributions of  $\chi^2_{H_0} - \chi^2_{H_1}$  in the reference case and for various values of  $\alpha_{qnl}$  will be produced and  
 3512 studied in Section 7.7.

3513 The idea behind this joint fit is that by letting the oscillation parameters and  $\delta \sin^2(2\theta_{12})$  and  $\delta \Delta m_{21}^2$   
 3514 free to float, converging potentially to arbitrary, wrong values in the case of oscillation parameters,  
 3515 we add some flexibility to fully exploit the difference introduced by unexpected effects between the  
 3516 reference spectra and correlations.

3517 There were other reasons to develop this joint fit. The main one was that it required an update of  
 3518 our software framework so it's able to perform joint fit. It was not fully ready for that. This feature

$\sin^2(2\theta_{12})$	$\Delta m_{21}^2$	$\Delta m_{31}^2$	$\sin^2(2\theta_{13})$
$0.851^{+0.020}_{-0.018}$	$7.53 \pm 0.18 \times 10^{-5} \text{ eV}^2$	$2.5283 \pm 0.034 \times 10^{-3} \text{ eV}^2$	$0.08523 \pm 0.00268$

TABLE 7.3 – Nominal PDG2020 value [92]. All value are reported assuming Normal Ordering.

3519 will be very useful when the Subatech team will include the TAO spectrum (via a joint fit) in the  
 3520 oscillation studies it will perform.

### 3521 7.2.5 Joint fit of the SPMT and LPMT spectra : distribution of $\delta \sin^2(2\theta_{12})$ and 3522 $\delta \Delta m_{21}^2$

3523 The last test statistics we will study might be complementary to  $\Delta \chi^2 = \chi^2_{joint,H0} - \chi^2_{joint,H1}$ .

3524 These test statistics are simply the fitted values of  $\delta \sin^2(2\theta_{12})$  and  $\delta \Delta m_{21}^2$ . In the reference case,  
 3525 when no unexpected reconstruction problem is present, we expect them to be distributed in an  
 3526 approximate gaussian way, centered on 0. When QNL effect will be included, they will tend to  
 3527 converge to higher values, to compensate the bias introduced on the fitted  $\sin^2(2\theta_{12})$  and  $\Delta m_{21}^2$  due  
 3528 to the distortion of the LPMT spectra and of the correlations between the LPMT and SPMT spectra.

3529 Again, these distributions will be studied in Section 7.7.

### 3530 7.2.6 Limitations

#### 3531 QNL in backgrounds

3532 The JUNO commons inputs provides background spectra that already have been smeared by the  
 3533 LPMT resolution. Because the resolution depends on  $E_{vis}$  (Eq. 7.19), to apply supplementary QNL we  
 3534 would need to de-convolute the LPMT resolution, apply the supplementary QNL then re-smear the  
 3535 spectra. This deconvolution is no trivial. Thus we ignore the background when produced distorted  
 3536 spectra.

3537 This should not affect too much the power of our statistical tools, as the backgrounds are common to  
 3538 both spectra and should not have any effect on the statistical covariance matrix.

#### 3539 Systematics

3540 It would be more rigorous to also include systematic uncertainties. However, in the present state of  
 3541 our fit framework, it would require the computation (often empirical, via the generation of thousands  
 3542 of toy samples) of  $(820 \times 820)$  covariance matrices, which was judge too time consuming with respect  
 3543 to the time we could devote to this chapter.

3544 Moreover, it seems reasonable to think that the sensibilities evaluated with only statistical uncer-  
 3545 tainties would not be changed much by a full treatment. Indeed, all the systematic uncertainties  
 3546 affect the true visible energy spectrum, before reconstruction. This spectrum is a common input to  
 3547 both the LPMT and SPMT reconstructions. Therefore, observed differences between the oscillation  
 3548 parameters measured by one or the other system should not be due to these systematics effects, and  
 3549 remain of the same order as if these effects were absent.

---

### 3550 Correlation between LPMT and SPMT reconstruction

3551 Most of our results assume uncorrelated reconstruction uncertainties between the SPMT and LPMT  
 3552 systems. In practice, once the  $E^{vis}$  of a toy event is generated (see Section 7.3.3), we simulate the  
 3553 SPMT and LPMT reconstruction by adding a  $\delta E_{SPMT}^{rec}$  and a  $\delta E_{LPMT}^{rec}$ .

3554 The two latter increments are chosen randomly on Gaussian distribution. These two drawings are  
 3555 carried out independently. In reality, the reconstruction of  $E^{vis}$  is about proportional to the number  
 3556 of PE, therefore to the number of scintillation photons produced in the scintillator. Both the LPMT  
 3557 and SPMT reconstruction depend on the stochastic variation of this number event to event. Their  
 3558 results therefore vary in a correlated way. The correlation is kept low since it is shuffled by another  
 3559 source of variability, namely the sampling of photons : the SPMT indeed reconstruct only a few  
 3560 dozen PEs when more than 10000 photons are emitted.

3561 This correlation is higher when the interaction takes place close to the sphere's surface (ie close to  
 3562 some of the PMTs), the non-uniformity effect is correlated between the two systems. To account  
 3563 for it, when should ideal produce the simulated samples necessary to our studies by using the full  
 3564 simulation. However, it would be far too CPU intensive. The impact of neglecting this correlation  
 3565 will be discussed in Section 7.5.

### 3566 Realistic QNL

3567 The way we implement the QNL effect in toy samples is also simplified. The size of the QNL effect  
 3568 in a PMT depends on the number of photons hitting it, therefore on the position of the interaction.  
 3569 When generating toy events, we apply QNL event-wise, only as a function of the value of  $E^{vis}$  (Eq.  
 3570 7.2). As explained in Section 7.1.2, the full simulation has been used to find the average  $\alpha_{qnl}$  for a  
 3571 given  $\gamma_{qnl}$  which is considered sufficient for this exploration.

3572 Again, replacing toy samples with samples generated with the full simulation would yield more  
 3573 accurate results, but is prohibitive in terms of calculation time. For future studies, sophisticated  
 3574 solutions to this problem will have to be found, but are out of the scope of this thesis.

## 3575 7.3 Fit software

3576 In this section, I describe the fit framework that was used in this study. The AveNu<sub>e</sub> framework is  
 3577 the adaptation to JUNO of one of the frameworks, partly developed at Subatech, used by the Double  
 3578 Chooz [140] experiment. It is composed of two parts: the AveNu<sub>e</sub> Generators and the AveNu<sub>e</sub> Fitting  
 3579 Package. The Generators are a set of standalone macros, the Fitting Package is an C++ package, using  
 3580 the RooFit library.

3581 Both parts of the package are interfaced with what we call the JUNO inputs. These inputs comprise  
 3582 all the ingredients to build a  $T(\theta, \eta)$  prediction, among which :

- 3583 — Reactor antineutrino spectra for each isotope as predicted by Mueller [141].
- 3584 — The isotopes mean releases energy.
- 3585 — Reactors' thermal powers and fission fractions.
- 3586 — Various corrections to account for the contributions from the Non Equilibrium Regime and the  
 3587 Spent nuclear fuel.
- 3588 — A correction obtained by comparing these spectrum prediction in the case of the Daya Bay  
 3589 experiment with actual Daya Baya data [69].
- 3590 — The IBD differential cross section as function of the antineutrino energy.

- The assumed values of the oscillation and nuisance parameters at the start of the fit or for sensitivity studies.
- Parameters describing the non linearity of the photon emission as a function of the deposited energy.
- Energy reconstruction parameters (see equation 7.19 and Figure 7.10).
- The selected IBD and background expected yields per day, and the background spectra, all obtained from JUNO’s full simulation and studies to design the selection.
- Uncertainties on all these quantities for the computation of covariance matrices.

We describe in the next section the role of each part of the framework.

### 7.3.1 AveNu<sub>e</sub> Standalone Generators

The main macro here is the “IBD generator” macro. It is used to :

- Compute  $T_{no\ osc}(\eta)$  (unoscillated theoretical spectra) predictions. It is done by toy generating a spectrum. In order to not be affected by statistical fluctuations, it generates 100 times more statistics than JUNO’s expected yield after 6 years. It is provided in the form of a TTree. These predictions concern a non oscillated spectrum.
- Toy samples simulated data sets. It is essentially used to simulate data spectra altered by QNL effects (see below).
- The above productions are input to the Fitting Package, or to other macros from Standalone Generators, which compute the covariance matrices necessary to the Fitting Package. Some of the covariance matrices are computed from the  $T$ ’s, using linear error propagation, some other are empirical calculations based on sets of toy samples generated with varying parameters. This is also the case for one of the versions of the computation of the  $V_{stat}$  covariance matrix of the LPMT+SPMT double spectrum (see Section 7.5).

### 7.3.2 AveNu<sub>e</sub> Fitting Package

Its role is to perform fits to a single data samples, or to a set of toy samples. In practice :

- It loads TTrees containing the data to fit as well as the  $T_{no\ osc}(\eta)$  predictions, and create local objects representing the data spectrum and the pdf. For that purpose,  $T_{no\ osc}(\eta)$  are changed into predictions  $T(\theta, \eta)$  for the oscillated spectrum by weighting events in the TTree according to the oscillation probability.
- It loads the necessary covariance matrices.
- It creates from this a  $\chi^2$  object. The Pearson, Neyman, CNP and Pearson V versions are available
- It is interfaced with Minuit via RooFit classes to perform the minimisation. At each step,  $T(\theta, \eta)$  are re-weighted by the oscillation probability corresponding to the current value of the floating oscillation parameters.

Three kinds of data can be fitted with this Package : real data, Asimov simulated data and toy data.

When real data will be available at JUNO, we expect that the result of the IBD selection will be made available by the collaborations via TTrees.

3631 The principles of Asimov fits were described in Section 2.7.3. In practice, our Fit Package fill the  
 3632 local object representing the data spectrum with  $T(\theta, \eta)$ , assuming some values for the oscillation  
 3633 parameters.

3634 The toy data samples can have two origins. Some are produced by the IBD generator macro of the  
 3635 AveNue Generators. This is the case of the toy samples that we produce with QNL effects. It is  
 3636 also possible to generate toys directly with the Fitting Package. In that case, toy data spectra are  
 3637 produced by generating random fluctuations around each the values of  $T(\theta, \eta)$ . These fluctuations  
 3638 must be the reflect of both statistical and systematic uncertainties. Fluctuations between bins  $i$  and  $j$   
 3639 can be correlated. Such correlations are common in the case of systematic uncertainties. In general,  
 3640 they are 0 for the statistical uncertainties. In our case, as already explained earlier (see for instance  
 3641 the Sections where the test statistics are described), bins from the SPMT part of the LPMT+SPMT  
 3642 spectrum are correlated to bins of the LPMT part even for the statistical part.

3643 To generate correlated fluctuation we use, through Choleski decomposition, the covariance matrices.  
 3644 This way to generate toy is faster. We use it in this work in the reference case (no QNL). In the case  
 3645 where QNL effects are simulated, the corresponding statistical covariance matrix is not known, we  
 3646 therefore resort to the IBD generator.

### 3647 7.3.3 Details of the IBD generator

3648 The IBD generator is a standalone generator used to produce oscillated and non oscillated spectra  
 3649 as the one seen by the JUNO experiment. It is at the core of the fitting framework as it's used to  
 3650 generate  $T(\theta, \eta)$ , the toy data and spectra to compute the covariances matrix.

3651 With thus have a flexible macro with options allow to enable or disable effects such as non-uniformity  
 3652 and non-linearity. It take as an argument the number of events to generate  $N_{evt}$ . Optionally, we  
 3653 generate an effective number of events  $N$  by drawing in a Poisson distribution of mean  $N_{evt}$ .

3654 Then for each event we:

- 3655 1. Choose randomly, following the reactor power fraction, the source reactor of the neutrino.
- 3656 2. Generate a random interaction position in the detector following a uniform distribution over  
 3657 the detector volume.
- 3658 3. Draw a random neutrino energy  $E_\nu$  from the expected neutrino emission spectrum of every  
 3659 reactor. This spectrum is computed by:
  - 3660 (a) Computing the power spectrum of each isotopes  $^{235}\text{U}$ ,  $^{238}\text{U}$ ,  $^{239}\text{Pu}$ ,  $^{241}\text{Pu}$  using the Huber-  
 3661 Mueller model [63, 66].
  - 3662 (b) Summing the contribution of each isotopes following the respective fission fraction [0.58,  
 3663 0.07, 0.30, 0.05] as reported in [142].
  - 3664 (c) The power of each reactor is then adjusted by their distances from the detector, the detec-  
 3665 tor efficiency and their mean duty cycle (11 of 12 month).
  - 3666 (d) The total spectrum is then finally adjusted by taking into account the correction of the Day  
 3667 Bay bump [69], adjustment due to spent nuclear fuel and due to the non-equilibrium.
- 3668 4. (Optional) Compute the survival probability due to oscillation at nominal oscillation parameters  
 3669 value. If the neutrino does not survive, the event is rejected and the algorithm restart from step  
 3670 (1).
- 3671 5. Compute the emitted positron energy  $E_{pos}$  from the mass difference. If the neutrino does not  
 3672 have enough energy reject the event and start from step (1).
- 3673 6. Compute the deposited energy  $E_{dep}$  by incrementing  $E_{pos}$  by 511 keV to account for the positron  
 3674 annihilation. We do not consider cases where some of the energy leak outside of the detector  
 3675 (positron or annihilation gammas escaping the CD).

- 3676 7. Correct the deposited energy with the expected event-wise non-linearity from [84] to obtain  
 3677 the visible energy  $E_{vis}$ .
- 3678 8. (Optional) Add a custom non-linearity as described in Section 7.1.2. This non linearity is char-  
 3679 acterized by  $\alpha_{qnl}$  to obtain  $E_\alpha$ .
- 3680 9. Finally, using the expected resolution of the LPMT and SPMT systems, provided in the JUNO  
 3681 common inputs, we draw from a gaussian characterized by those resolution the reconstructed  
 3682 energy  $E_{rec}$  or  $E_{lpmt}$  and  $E_{spmt}$  for each systems. The resolutions are provided as ABC parame-  
 3683 ters using

$$\frac{\sigma E_{vis}}{E_{vis}} = \sqrt{\left(\frac{A}{\sqrt{E_{vis}}}\right)^2 + B^2 + \left(\frac{C}{E_{vis}}\right)^2} \quad (7.19)$$

3684 where A is the term driven by the Poisson statistics of the total number of detected photoelec-  
 3685 trons, C is dominated by the PMT dark noise, and B is dominated by the detector's spatial  
 3686 non-uniformity. The relative and absolute resolutions of the LPMT and SPMT systems are  
 3687 illustrated in Figure 7.10.

3688 The events are stored as n-tuples and are not yet binned at the end of the generator.

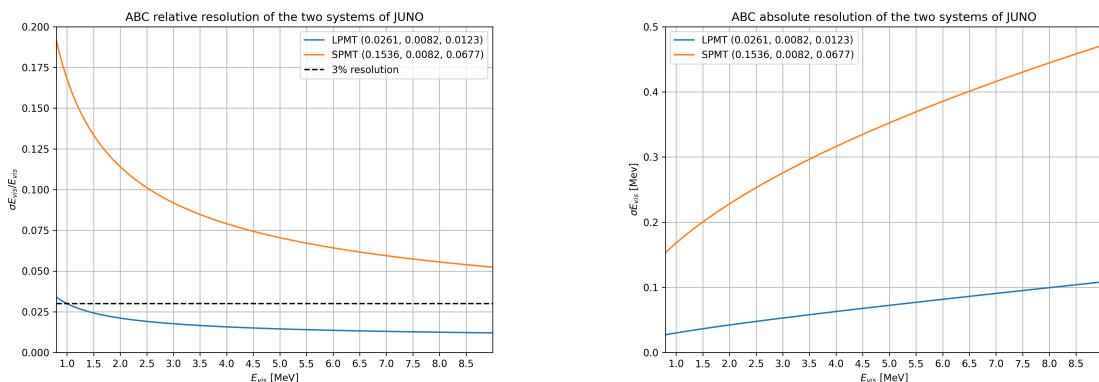


FIGURE 7.10 – Relative (On the left) and absolute (On the right) resolutions of the LPMT and SPMT systems used in this study. The number in parenthesis are the parameter A, B and C respectively for each systems.

## 3689 7.4 Technical challenges and development

3690 The fit framework Avenue was already partially developed with multispectra fitting in mind but  
 3691 a lot technical development was necessary to allow for a joint fit. This required a lot of work and  
 3692 constitute a good part of my total effort on this study. I remind that these development will be useful  
 3693 beyond this thesis and this subject. As already mentioned earlier, at some point, we should perform  
 3694 simultaneous fits of the JUNO and TAO spectra. It's also a potential starting point for combined  
 3695 analyses with other experiments, like long baseline experiments.

3696 The first step was to migrate the framework from ROOT5 (last release in March 2018) to ROOT6  
 3697 (v6.26.06 released in July 2022) to ensure compatibility with the data coming from the JUNO collab-  
 3698 oration, and benefiting of the improvement and corrections that came with ROOT6. This allow us to  
 3699 upgrade the C++ standard from C++11 to C++17. A substantial effort has been done to modernize  
 3700 the code, generalizing the functions and methods via templating to help readability and using smart  
 3701 pointer to prevent possible memory leaks.

3702 The Avenue framework had to be adapted, notably on the chi-square calculation and spectrum  
 3703 generation to correctly take into account the correlation between the SPMT and LPMT spectra.  
 3704 The delta joint fit requiring two more parameters over a spectrum twice as large as before with  
 3705 LPMT takes much more time, around 15h for 6 years exposure, than the single LPMT fit. Thus the  
 3706 framework and the fit macro had to be updated for distributed computing. Notably the aggregation  
 3707 of fit results can now be done in a single file instead of managing a file per fit. In case of numerous  
 3708 toy, the hard drive access time could lead to long analysis time.

3709 While the IBD generator was already able to generate LPMT and SPMT spectrum, it was not designed  
 3710 for generating correlated spectrum. As detailed in Section 7.3.3, up to the reconstruction effect, the  
 3711 two spectrum need to share the same generation else the two spectrum would be decorrelated and it  
 3712 would be like we would run two different experiment.

## 3713 7.5 Covariance matrix

3714 The covariance matrix between the LPMT and SPMT spectra is at the heart of this study as it was  
 3715 already mentioned in Section 7.2. In this section we discuss the different approaches taken to estimate  
 3716 it. We remind that in this work, we consider only statistical effects and let to future works the  
 3717 task to include systematic uncertainties. We thus evaluate in this section the  $(820 \times 820)$  statistical  
 3718 covariance matrix  $V$  of the LMPT+SPMT spectrum.

3719 As already explained in previous Sections 7.2.6 and 7.3.3, we assume, in most of what follows, that  
 3720 the effect of the energy reconstruction resolution is independent between the LPMT system and the  
 3721 SPMT system, although this is an approximation. We therefore also briefly study the correlations  
 3722 between the two reconstructions.

### 3723 7.5.1 Analytical method

3724 The first method discussed is the analytical method where we propagate the resolution of the LPMT  
 3725 and SPMT spectra over a non-smeared spectrum. Following the approach used in the IBD generation  
 3726 in Section 7.3.3, we consider the system resolution  $\sigma(E)$  to be only dependent in energy. We do not  
 3727 consider the position of the event.

3728 Using the formalism of section 39.2.5 *Propagation of errors* of PDG2020 [92] and considering an ex-  
 3729 tended spectrum of 820 bins following the binning scheme introduced in 2.7.2, the first 410 for the  
 3730 LPMT and the last 410 for the SPMT, we consider

- 3731 —  $\mathbf{h} = (h_0, \dots, h_n)$  Is the n-dimensional vector ( $n=820$ ) containing the number of entries in each  
 3732 bin of the LPMT+SPMT true  $E^{vis}$  spectrum.
- 3733 —  $\zeta(\mathbf{h}) = (\zeta_0(\mathbf{h}), \dots, \zeta_n(\mathbf{h}))$  is the n dimensional vector containing the reconstructed  $E^{vis}$  LPMT+SPMT  
 3734 spectrum.

3735 Since, like in most sensitivities studies, resolution is simulated via a gaussian smearing,  $\zeta$  can be  
 3736 expressed this way :

$$\zeta_i = \sum_{j=0}^n G(j, \sigma(E_j))(i) \cdot h_j \quad (7.20)$$

3737 where  $G(j, \sigma(E_j))(i)$  is the smearing function defined as

$$G(j, \sigma(E_j))(i) = \int_{\lfloor E_i \rfloor}^{\lceil E_i \rceil} \frac{1}{\sigma(E_j) \sqrt{2\pi}} e^{-\frac{(E_j - E)^2}{2\sigma(E_j)^2}} dE \quad (7.21)$$

3738 where  $E_j$  is the mean energy in the bin  $j$  and  $\lfloor E_i \rfloor$  and  $\lceil E_i \rceil$  are the lower and higher energy bound of  
 3739 the  $j$ th bin respectively.

3740 According to 7.21, to evaluate  $V$ , the matrix describing the uncertainties on  $\zeta_i$ 's and the correlations  
 3741 between them, one has to consider uncertainties both on  $h_j$ 's and on  $G(j, \sigma(E_j))(i)$ 's. We use linear  
 3742 error propagation and split this problem in two steps :  $V = V_{inputs} + V_{rec}$ . The first matrix accounts  
 3743 for the uncertainties on the inputs from the true  $E^{vis}$  spectrum ( $h_i$ 's), while the second concerns the  
 3744 uncertainties due to  $G(j, \sigma(E_j))(i)$ 's.

3745 To evaluate  $V_{inputs}$ , we use  $V_{inputs} = AUA^T$  where  $U$  is the covariance matrix of the LPMT+SPMT  
 3746 true  $E^{vis}$  spectrum. Since before reconstruction the LPMT and SPMT spectra are the same, this  
 3747 LPMT+SPMT is the juxtaposition of two 410-bin identical spectra. Moreover we are interested only  
 3748 in statistical uncertainties. Therefore,  $U$  has the form :

$$U = \begin{cases} \sqrt{h_i h_j} & \text{if } i = j \text{ or } |i - j| = 410 \\ 0 & \text{otherwise} \end{cases} \quad (7.22)$$

3749 The condition  $|i - j| = 410$  express the fact that one  $h_i$  of the LPMT part of the spectrum is naturally  
 3750 100% correlated with the corresponding bin in the SPMT spectrum.

3751 We can then construct the transfer matrix  $A$  as

$$A_{ij} = \frac{\partial \zeta_i}{\partial h_j} = G(j, \sigma(E_j))(i) \quad (7.23)$$

3752 and then compute the first part of our covariance matrix

$$V_{inputs} = AUA^T \quad (7.24)$$

3753 Now we need to consider the uncertainty on the steaming from the resolution, ie to evaluate  $V_{rec}$ . It  
 3754 can be done considering no uncertainty on the true  $E^{vis}$  spectrum. The quantity  $G(j, u) \equiv G(j, \sigma(E_j))(i)$   
 3755 is the predicted probability for an event initially in bin  $j$  of the true  $E^{vis}$  spectrum to be reconstructed  
 3756 in bin  $i$ . In practice, the migration between these bins is a random process. Reconstructed many times  
 3757 the same event would not lead each time the same migrations. We need here to determine this  
 3758 variability. We consider that with 410 bins, migrations vary independently whatever  $i$  and  $j$ .

3759 This allows to consider  $V_{rec}$  as diagonal, thus we only need  $\sigma G(j, i)$ . We can derive this term from  
 3760 two equation:

- 3761 — The term  $G(j, i) \cdot h_j$  represent the number of event smeared from the bin  $j$  that end up in the  
 3762 bin  $i$ . This is a number, we thus assume poissonian statistic so that  $\sigma[G(j, i) \cdot h_j] = \sqrt{G(j, i) \cdot h_j}$ .
- 3763 — Using basic error propagation we can say that  $\sigma^2[G(j, i) \cdot h_j] = h_j^2 \sigma^2 G(j, i) + G(j, i)^2 \sigma^2 h_j$ .

Equating the above equations, and remembering that  $\sigma h_j = \sqrt{h_j}$  since  $h_j$  is also a number of events :

$$G(j, i)h_j = \sigma^2[G(j, i)h_j] = h_j^2 \sigma^2 G(j, i) + G(j, i)^2 h_j \quad (7.25)$$

$$\Rightarrow \sigma^2 G(j, i) = \frac{G(j, i)h_j - G(j, i)^2 h_j}{h_j^2} \quad (7.26)$$

$$= \frac{(1 - G(j, i))G(j, i)}{h_j} \quad (7.27)$$

3764 By summing the two covariance matrix  $V_{inputs}$  and  $V_{rec}$ , we can extract a correlation matrix presented  
 3765 in Figure 7.11. Typically, a bin in the SPMT part of the reconstructed spectrum is correlated up to  
 3766 a few percents to the corresponding bin in the LPMT spectrum and its neighbour. This might seem  
 3767 a small correlation. However, its concerns all bins. The global impact is therefore high. As an  
 3768 illustration, as seen in Section 7.2.2, the correlation between the value of  $\sin^2(2\theta_{12})$  measured with  
 3769 the LPMT spectrum and that measured with the SPMT spectrum are correlated at more than 95%.

3770 The correlation between the SPMT and LPMT spectra is greater at the start of the spectrum. This  
 3771 is expected since the absolute resolution is smaller in this region. For instance, at 1.5 MeV, the  
 3772 reconstruction by the SPMT re-distribute events with a sigma of more than 0.20 MeV. At 6 MeV, this  
 3773 is about twice more. Since the resolution reduces the initial correlations (true Evis spectra are share  
 3774 by both LPMT and SPMT, correlations are 100%), we therefore expect higher remaining correlations  
 3775 where the absolute resolution is smaller.

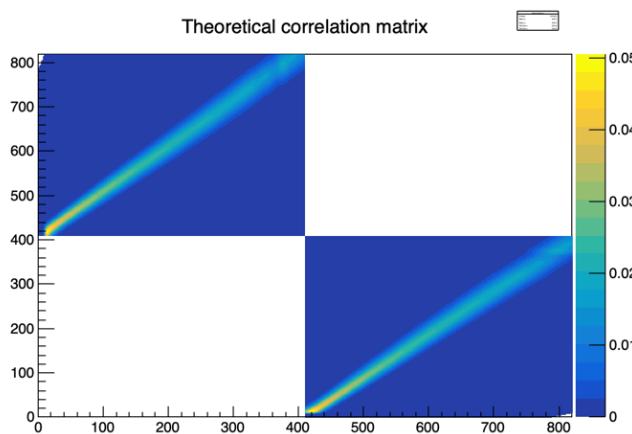


FIGURE 7.11 – Theoretical correlation matrix between the LPMT spectrum (bins 0-409) and the SPMT spectrum (410-819). The diagonal has been set to 0 (it was 1) for readability purpose.

### 3776 7.5.2 Empirical method

3777 The second method is the empirical way where we generate toys and just compute the empirical  
 3778 correlation between the bin contents.

$$\text{Corr}(h_i, h_j) = \frac{\mathbb{E}[h_i h_j] - \mathbb{E}[h_i] \mathbb{E}[h_j]}{\sigma_{h_i} \sigma_{h_j}} \quad (7.28)$$

3779 We thus generate  $10^7$  event using the IBD generator presented in Section 7.3.3, then produce spectra  
 3780 from this finite set of events, meaning we must choose a number  $N$  of toy each composed of  $M$  event  
 3781 in order to have the best estimate.

3782 It can be shown that empirical correlations are more precise when one maximises the number of  
 3783 samples, even at the price to have few events per sample. This effect is illustrated in Figure 7.12.

3784 The relative difference between the element of the theoretical matrix of Figure 7.11 and the empiric  
 3785 correlation matrix in Figure 7.12c is presented in Figure 7.13. Typically, correlations coefficient differ  
 3786 by 20% of their value. We have verified that differences larger than this are confined in the very low  
 3787 or high end of the energy spectrum, which carry no sensitivity to the solar oscillation parameters we

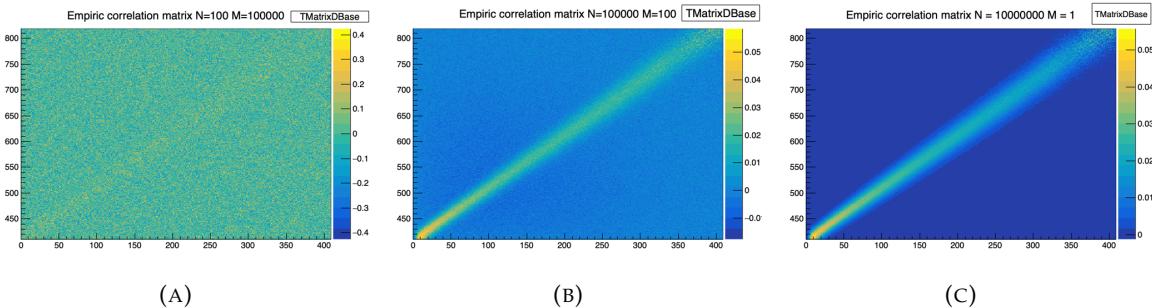


FIGURE 7.12 – Upper left corner of the estimated correlation matrix between the LPMT and SPMT spectrum for different configuration of  $N$  toy with different number of  $M$  events per toy. We observe that the statistical uncertainty, the noise effect, diminish with the number of toy considered.

3788 aim at. Therefore, for the statistical tests presented in this chapter we assume the correlations present  
3789 in the theoretical version of  $V$ . This should account for the effect of correlations well enough.

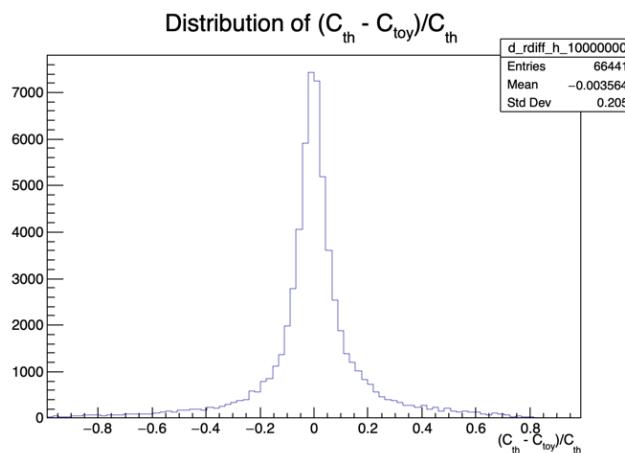


FIGURE 7.13 – Relative difference between the element of the theoretical and empiric correlation matrix

We chose to do so for practical reasons. Indeed, for the  $\chi^2$  computation and the Choleski decomposition (see Section 7.3.2) the matrix must be invertible and positive definite. The statistical uncertainty on the coefficient of the empiric matrix can prevent that, leading to complications.

3793 7.6 Technical Validation

3794 Standard Independent Joint Fit

We have already explained in Sections 7.2 and 7.5 that a correlation exist between the SPMT and LPMT spectra, and is accounted for in the LPMT+SPMT joint fit by the  $V$  covariance matrix, which determination is described in the previous section.

We can, however, perform a test where we ignore these correlations, setting to 0 all off-diagonal elements of  $V$ . In this case, we implicitly assume that our data contains more information, and therefore expect the uncertainties on  $\Delta m_{21}^2$  and  $\sin^2(2\theta_{12})$  to be smaller than those obtained with

3801 individual fits to the LPMT and SPMT spectra. Assuming a gaussian behavior of the number of  
 3802 entries per bin, these uncertainties should be close to the weighted average of the uncertainties with  
 3803 the individual fits :

$$\frac{1}{\sigma_{Weighted}^2} = \frac{1}{\sigma_{LPMT}^2} + \frac{1}{\sigma_{SPMT}^2} \quad (7.29)$$

3804 These tests are performed using an Asimov sample. Indeed, if it was done via a toy study, then  
 3805 generating correlated toy spectra and fitting them assuming a diagonal  $V$  matrix would have led to  
 3806 biases, regardless of the quality of the technical implementation. Asimov spectra, on the other hand,  
 3807 are generated with no fluctuations. They are supposed to return fitted values of  $\Delta m_{21}^2$  and  $\sin^2(2\theta_{12})$   
 3808 exactly equal to the values assumed during the generation. This is, together with the comparison  
 3809 with  $\sigma_{weight}$ , a strong test of the technical implementation.

3810 Note that we fix here the  $\delta m_{21}^2$  and  $\delta \sin^2(2\theta_{12})$  parameters to 0. Also, we assume 6 years of data  
 3811 taking, and the absence of unexpected instrumental effects (no supplementary QNL). A notable  
 3812 difference with the fit configuration used later in this chapter (and presented in Section 7.2) is that  
 3813 we do not treat  $\Delta m_{31}^2$  as a nuisance parameter. It is free to float.

	$\sigma(\Delta m_{21}^2)$ [eV $^2$ ]	$\sigma(\delta \Delta m_{21}^2)$ [eV $^2$ ]	$\sigma(\sin^2(2\theta_{12}))$	$\sigma(\delta \sin^2(2\theta_{12}))$	$\sigma(\Delta m_{31}^2)$ [eV $^2$ ]	$\chi^2$
LPMT	$1.29 \times 10^{-07}$		$1.33 \times 10^{-03}$		$4.39 \times 10^{-06}$	$3.23 \times 10^{-18}$
SPMT	$1.38 \times 10^{-07}$		$1.38 \times 10^{-03}$			$2.87 \times 10^{-18}$
Indep Standard joint	$9.48 \times 10^{-08}$		$9.86 \times 10^{-04}$		$4.39 \times 10^{-06}$	$6.10 \times 10^{-18}$
Standard joint	$1.29 \times 10^{-07}$		$1.18 \times 10^{-03}$		$4.39 \times 10^{-06}$	$3.38 \times 10^{-18}$
Weighted	$9.46 \times 10^{-08}$		$9.63 \times 10^{-04}$			
Delta joint	$1.35 \times 10^{-07}$	$3.43 \times 10^{-08}$	$1.38 \times 10^{-03}$	$1.46 \times 10^{-04}$	$4.39 \times 10^{-06}$	$3.38 \times 10^{-18}$
Indep Delta joint	$1.38 \times 10^{-07}$	$1.89 \times 10^{-07}$	$1.38 \times 10^{-03}$	$1.87 \times 10^{-03}$	$4.39 \times 10^{-06}$	$6.10 \times 10^{-18}$

TABLE 7.4 – Uncertainties on each parameters reported by Minuit on Asimov studies.  
 LPMT and SPMT rows are the results on the individual fit on each spectra. The Weighted row correspond to the weighted average uncertainties between the LPMT and SPMT fits following Eq. 7.29. The Indep Standard joint row is the result of the joint LPMT+SPMT fit but the off-diagonal terms are set to 0. The Indep Standard joint and Standard joint fits both are LPMT+SPMT fit but the parameters  $\delta m_{21}^2$  and  $\delta \sin^2(2\theta_{12})$  are fixed to 0. The Delta joint and Indep Delta joint are LPMT+SPMT fit with  $\delta m_{21}^2$  and  $\delta \sin^2(2\theta_{12})$ , difference being that in the Indep version, the off-diagonal terms of the covariance matrix are set to 0.

3814 The results are reported in Table 7.4. All those test are ran considering statistics error only, 6 years  
 3815 exposure with all backgrounds,  $\sin^2(2\theta_{13})$  fixed to its nominal value. For the SPMT individual  
 3816 fit  $\Delta m_{31}^2$  is fixed at its nominal value as the SPMT system is not sensitive to this parameter. We  
 3817 use here the simple Pearson  $\chi^2$ . Indeed, as explained above, an Asimov fit is supposed to find  
 3818 exactly the values of the parameters assumed for the generation of the spectrum, which implies a  
 3819 very low Pearson  $\chi^2$  (0 modulo numerical effects). This is also a strong indication that the technical  
 3820 implementation is correct. If we had used the usual Pearson V  $\chi^2$ , the  $\ln |V|$  term would have made  
 3821 the result more difficult to interpret.

3822 When we performed the Standard Independent Joint Fit, as expected we observed that the fitted  
 3823 values of the parameters all matched the generation values. We can also see in table 7.4 that the  
 3824 uncertainty on  $\Delta m_{21}^2$  evaluated by the fit are equals the corresponding  $\sigma_{weight}$  up to 0.2%. In the  
 3825 case of  $\sin^2(2\theta_{12})$ , the agreement is up to 2.5%.

3826 A slight difference exists in statistic between the SPMT and LPMT spectra. Indeed, due to a larger  
 3827 smearing in energy resolution, events that would be inside the spectrum range [0.8, 7.5] MeV are  
 3828 smeared outside it. The  $\sin^2(2\theta_{12})$  parameter being mainly driven by the amplitude of the spectrum  
 3829 (see illustration 7.1), it is more affected than  $\Delta m_{21}^2$ .

3830 **Standard Joint Fit**

3831 This case is similar to the previous one, with one difference : we now use the version of  $V$  that  
 3832 accounts for the correlations between the SPMT and LPMT spectra. The expected effect of this  
 3833 correlation is that the uncertainties on  $\Delta m_{21}^2$  and  $\sin^2(2\theta_{12})$  should see very little improvement with  
 3834 respect to individual fits.

3835 Moreover, the uncertainty on  $\Delta m_{31}^2$  should be very close to that obtained by the individual fit to  
 3836 the LPMT spectrum since only this one contains information on  $\Delta m_{31}^2$  (thanks to its high energy  
 3837 resolution). This is therefore a rather robust test.

3838 As can be seen in Table 7.4, these expectations are observed in practice.

3839 **Delta Joint Fit**

3840 It is the same fit as above, where we let the  $\delta \Delta m_{21}^2$  and  $\delta \sin^2(2\theta_{12})$  parameters free to float in the fit.  
 3841 A test assumes no correlations (diagonal  $V$ ), the other one assumes the usual  $V$ .

3842 A first test here is that the fitter should find these parameters at 0, since no QNL is introduced in these  
 3843 Asimov spectra. Also, in the correlated case, we expect the uncertainties on  $\delta \Delta m_{21}^2$  and  $\delta \sin^2(2\theta_{12})$   
 3844 to be far smaller than in the independent case. Indeed, when the  $\chi^2$  considers these two spectra are  
 3845 correlated, distorting only the LPMT part of the PDF without changing the SPMT part (remember:  
 3846  $\delta \Delta m_{21}^2$  and  $\delta \sin^2(2\theta_{12})$  appear only in  $T(\theta, \eta)$  for the 410 first bins, see Section 7.2) leads to a quick  
 3847 explosion of this  $\chi^2$  when profiling values of  $\delta \Delta m_{21}^2$  and  $\delta \sin^2(2\theta_{12})$  away from 0.

3848 Results in Table 7.4 are again consistent with these expectations.

3849 **Toy studies**

3850 The same tests as above have been repeated, using a set of 1000 toy samples instead of one Asimov  
 3851 sample. Only cases where we account for the correlations between the SPMT and LPMT spectra are  
 3852 carried out. The generation of the toy samples includes these correlations. We therefore also test that  
 3853 part.

3854 We can see on Figures 7.14 and 7.15 the distribution of the best fit values for all the parameters of  
 3855 interest. The mean values and standard deviations are in all cases consistent with the results of the  
 3856 Asimov tests (Table 7.4). Therefore, when realistic fluctuations are simulated, even with a peculiar  
 3857  $\chi^2$  computed with a complex covariance matrix and correlated data, the fit is stable and unbiased.

3858 These distributions also confirm that the uncertainties on  $\delta \Delta m_{21}^2$  and  $\delta \sin^2(2\theta_{12})$  are an order of  
 3859 magnitude smaller than the uncertainties on  $\Delta m_{21}^2$  and  $\sin^2(2\theta_{12})$ . This is an indication of the power  
 3860 of the test statistics used in this chapter.

3861 **Conclusion of the technical validation**

3862 All the tests carried out in this section are consistent with our expectation. We therefore conclude  
 3863 that the technical implementation of the tools used in this chapter is correct.

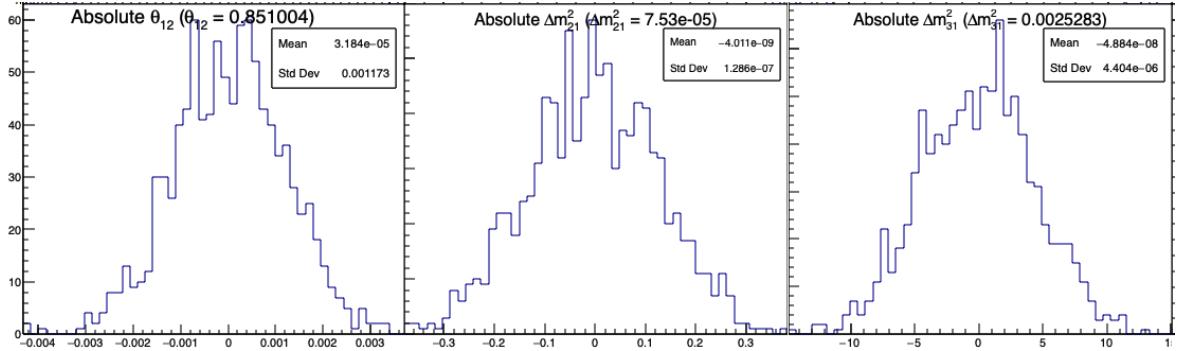


FIGURE 7.14 – Distribution of BFP - nominal value for 1000 toy Standard joint fit. 6 years exposure, all background, PearsonV  $\chi^2$ ,  $\theta_{13}$  fixed. In those plots,  $\theta_{12}$  stands for  $\sin^2(2\theta_{12})$

## 3864 7.7 Results

### 3865 7.7.1 Effect of supplementary QNL on the LPMT spectrum

3866 In this first part of this Section 7.7, we will present the sensitivity of various test statistics to un-  
 3867 expected instrumental effects affecting the SPMT and LPMT differently. The latter effects will be  
 3868 illustrated by generated toy samples affected by the QNL effect.

3869 Most of the tests involve either an individual fit to the LPMT spectrum or the SPMT spectrum, or  
 3870 a joint fit of these two spectra. To better understand why some test statistics turn out to be more  
 3871 powerful than others, we study briefly in the present subsection the results of these fits and interpret  
 3872 the differences.

3873 We generate toy spectra, and fit them according to the default configuration described in Section  
 3874 7.2. During the generation of the LPMT spectrum, we distort it to simulate a QNL effect, with an  
 3875 intensity of  $\alpha_{qnl} = 1\%$ . For reference, this is about three times the expected residual QNL after the  
 3876 application of dual calorimetric calibration methods ( $\alpha_{qnl} = 0.3\%$  [84]).

3877 Backgrounds had to be ignored here: the JUNO inputs described in Section 7.3 provide a recon-  
 3878 structed spectrum, but not the event per event information about the true  $E_{vis}$ , which we need to  
 3879 apply the QNL effect (See Equation 7.19).

3880 The effect of this QNL on the spectrum is illustrated in Figure 7.16 In Table 7.5 we report the results  
 3881 of the different kinds of fits.

3882 We notice (1st line, first 3 columns) that the individual fit to the LPMT spectrum tends to find, as  
 3883 expected, biased value for  $\Delta m^2_{21}$  and  $\sin^2(2\theta_{12})$  and  $\Delta m^2_{31}$  (biased at about -1 sigma, -1.3 sigma and  
 3884 - 2.2 sigmas respectively). When a joint fit is performed, with the  $\delta\Delta m^2_{21}$  and  $\delta\sin^2(2\theta_{12})$  fixed at 0,  
 3885 and ignoring in the computation of the  $\chi^2$  the correlations between the LPMT and SPMT spectra, the  
 3886 biases on  $\Delta m^2_{21}$  and  $\sin^2(2\theta_{12})$  (3rd line, first 3 columns) appear to be average of the biases seen by the  
 3887 individual fits to these spectra, a logical result since the individual sensitivities to these parameters  
 3888 are similar. The bias on  $\Delta m^2_{31}$  (3rd column) remains the same as with the individual fit to the LPMT  
 3889 spectrum, however, which is expected since the SPMT spectrum carries no sensitivity to  $\Delta m^2_{31}$ .

3890 When the joint fit is performed with the nominal covariance matrix (determined in Section 7.5 as-  
 3891 suming no QNL), biases on  $\Delta m^2_{21}$  and  $\sin^2(2\theta_{12})$  explode: they are, respectively, about 6.5 and 2.5  
 3892 times larger (4th line).

3893 We explain it by the following mechanism : the fit tries to improve the agreement between the PDF  
 3894 and the data in the LPMT part of the spectrum by choosing biased values of the parameters. This

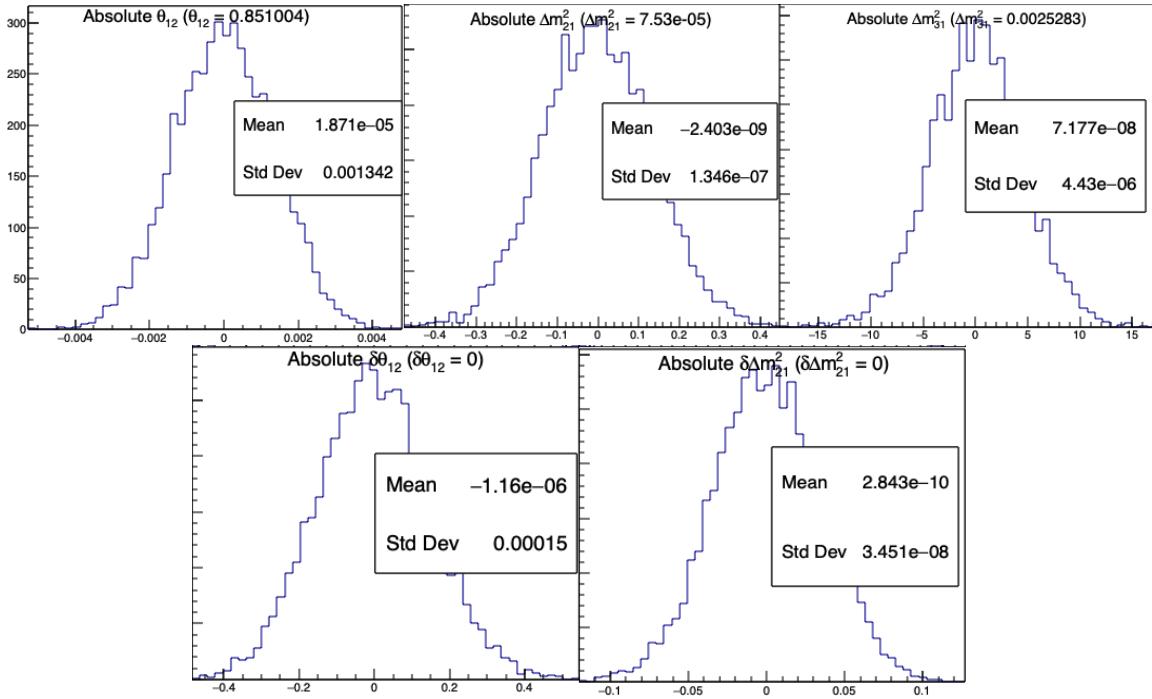


FIGURE 7.15 – Distribution of BFP - nominal value for 5000 toy Delta joint fit. 6 years exposure, all background, PearsonV  $\chi^2$ ,  $\theta_{13}$  fixed. In those plots,  $\theta_{12}$  stands for  $\sin^2(2\theta_{12})$  and  $\delta\theta_{12}$  for  $\delta \sin^2(\theta_{12})$

in turn tends to deteriorate the agreement between the PDF and the SPMT spectrum (not distorted by QNL). In the end, a discrepancy remains between data and PDF in at least one sector (LPMT or SPMT) if not both. When the  $\chi^2$  is built with a matrix which accounts for the correlations, this discrepancy can make the  $\chi^2$  explode.

For instance, in some bins of the LPMT spectrum, we can imagine the PDF overestimates the QNL-distorted data, while the contrary happens in the corresponding bins of the SPMT spectrum. If the expected correlation is positive between these two bins, the  $\chi^2$  will reach values accounting for a larger discrepancy than if no correlation existed and if only the raw agreement between the pdf and the spectra was important.

In reality, the consistency between the two can be judged only accounting for the correlations. This is the important role of the covariance matrix in this work. In other words, the spectra predictions are not only the  $T(\theta, \eta)$ 's, but also the correlations.

Another point must be noted : the correlation matrix  $V$  is evaluated assuming no QNL. With the QNL effect added, the actual correlations between the LPMT and SPMT generated toy spectra is a bit different, adding another source of discrepancy between the data and the predictions, and further increasing the  $\chi^2$ .

All in all, the minimisation of the  $\chi^2$  requires a larger scan of the oscillation parameters values than when correlations are ignored. Values can be chosen which are farther from the nominal ones, meaning larger biases.

This is actually an advantage. Indeed, we can see in table 7.5 that when  $\delta\Delta m^2_{21}$  and  $\delta \sin^2(2\theta_{12})$  are allowed to float in the fit, they "absorb" a large part of the bias. Notice in particular that adding the value of  $\delta\Delta m^2_{21}$  to the remaining bias on  $\Delta m^2_{21}$  (last line, columns 1 and 4) one retrieves the bias of the individual fit to the LPMT spectrum. The same applies to  $\sin^2(2\theta_{12})$ . Consequently, large values of  $\delta\Delta m^2_{21}$  and  $\delta \sin^2(2\theta_{12})$  are expected, hence high significances to help us to detect the distortion.

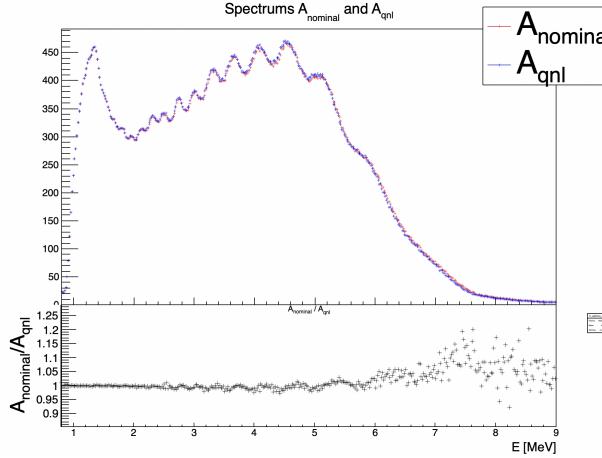


FIGURE 7.16 – **Top:** Theoretical spectrum without QNL (in red) and with  $\alpha_{\text{qnl}} = 1\%$  (in blue). **Bottom:** Ratio between the theoretical spectrum with and without QNL.

Mean (std dev)	$\theta_{12} [10^{-3}]$	$\Delta m_{21}^2 [10^{-7}\text{eV}^2]$	$\Delta m_{31}^2 [10^{-6}\text{eV}^2]$	$\delta\theta_{12} [10^{-3}]$	$\delta\Delta m_{21}^2 [10^{-7}\text{eV}^2]$
LPMT	-1.569 (1.171)	-0.957 (0.989)	-8.235 (3.898)	Irrelevant	Irrelevant
SPMT	-0.164 (1.191)	-0.603 (1.054)	Not sensitive	Irrelevant	Irrelevant
Indep Standard	-0.880 (1.174)	-0.786 (1.004)	-8.195 (3.900)	Irrelevant	Irrelevant
Standard	-8.106 (1.423)	-2.483 (1.018)	-6.649 (4.008)	Irrelevant	Irrelevant
Indep Delta	-0.169 (1.190)	-0.598 (1.054)	-8.234 (3.899)	-1.397 (0.259)	-0.361 (0.366)
Delta	-0.163 (1.183)	-1.532 (1.036)	-8.193 (3.934)	-1.441 (0.193)	0.654 (0.303)

TABLE 7.5 – In each column, the mean of the distribution of the 1000 best fit values found by fitting the 1000 toy samples with  $\alpha_{\text{qnl}} = 1\%$  is shown, from which we subtracted the value assumed when generating the toys. A value different from 0 indicates a bias. Between bracket, the average uncertainty of the fitted value is also shown. It allows to judge of the severity of the bias. For instance, the measurement of  $\sin^2(2\theta_{12})$  by fitting only the LPMT spectrum tends to be biased at the  $-1.569/1.171 = -1.34$  sigma.

In this case (last line, column 4 and 5), we see the most probable values of the fitted  $\delta\Delta m_{21}^2$  and  $\delta\sin^2(2\theta_{12})$  parameters differ from zero at about 7.46 sigma and 2.2 sigma.

Based on the above observations, we expect the " $\chi_{H_0}^2 - \chi_{H_1}^2$ " and "Distributions of  $\delta\Delta m_{21}^2$  and  $\delta\sin^2(2\theta_{12})$ " test statistics described in sections 7.2.4 and 7.2.5 to have the highest power. The "Direct comparison between the SPMT and LPMT spectra" should perform in the same ballpark. Finally, the "Comparison of individual fits" is expected to be have less power.

### 7.7.2 Comparison and statistical tests results

I present in this following Subsection the results from the tests and comparison detailed in section 7.2. For each distribution we compute the median p-value with respect to the distribution  $\mathcal{D}(\alpha_{\text{qnl}} = 0\%)$ . For this, we compute the median value of the distribution of interest  $\mathcal{D}(\alpha_{\text{qnl}})$ , then compute the p value

$$p = \frac{N(\mathcal{D}(0) > \text{Median}[\mathcal{D}(\alpha_{\text{qnl}})])}{N_{\text{tot}}} \quad (7.30)$$

where  $N(\mathcal{D}(0) > \text{Median}[\mathcal{D}(\alpha_{\text{qnl}})])$  is the number of toy in the distribution  $\mathcal{D}(\alpha_{\text{qnl}} = 0\%)$  that have a greater value than the median of the  $\mathcal{D}(\alpha_{\text{qnl}})$ . The p-value represent the probability for a non perturbed event to do worse that the median perturbed event.

3933 The uncertainty on the p-value is computed using

$$\sigma p = \sqrt{\frac{p(1-p)}{N}} \quad (7.31)$$

3934 which do not account for all uncertainties but serves as indicator.

### 3935 Comparison of solar parameters from individual analysis: $\chi^2_{ind}$

3936 The results are presented in Figure 7.17. We see that the p-value are much less significant than the  
 3937 other tests, this is because this test possess much less information about the relation between the  
 3938 LPMT and SPMT systems.

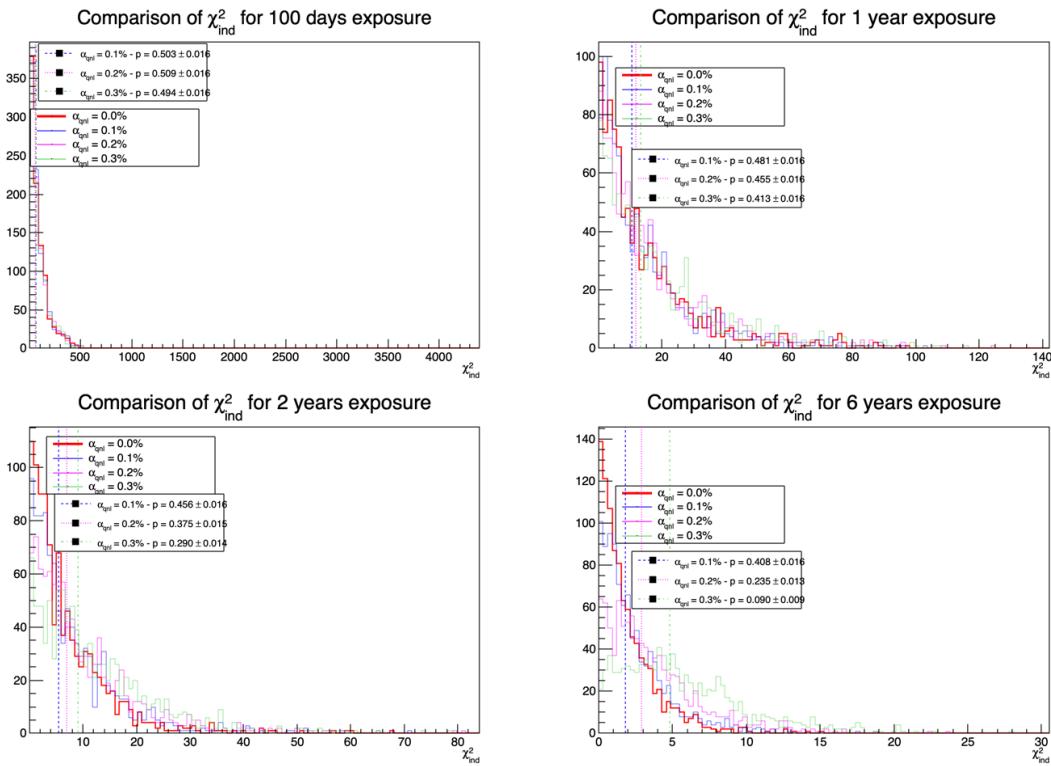


FIGURE 7.17 – Distribution of the  $\chi^2_{ind}$  for 1000 toys for different exposures. The dashed lines represent the median of the distributions and the p-value are the percentage of the  $\alpha_{qnl} = 0$  distribution that are greater than those medians.

3939 This test is the most straightforward as it require only the fit of the two spectra and the estimation of  
 3940 the parameters covariances, but is also the less powerful with a p value for  $\alpha_{qnl} = 0.3\%$  of  $0.09 \pm 0.009$   
 3941 at 6 years.

### 3942 Direct comparison between the LPMT and SPMT spectra: $\chi^2_{spe}$

3943 The results for different exposures can be found in Figure 7.18. To give an idea of the significance of  
 3944 this test, we provide the median p-value for each test  $\alpha_{qnl} \neq 0$ . As expected, the power of this test  
 3945 rises as the exposure does. We see significant discrimination at 6 years for  $\alpha_{qnl} \geq 0.3\%$  where the  
 3946 p-value for  $\alpha_{qnl} = 0.3\%$  is  $0.005 \pm 0.0022$ .

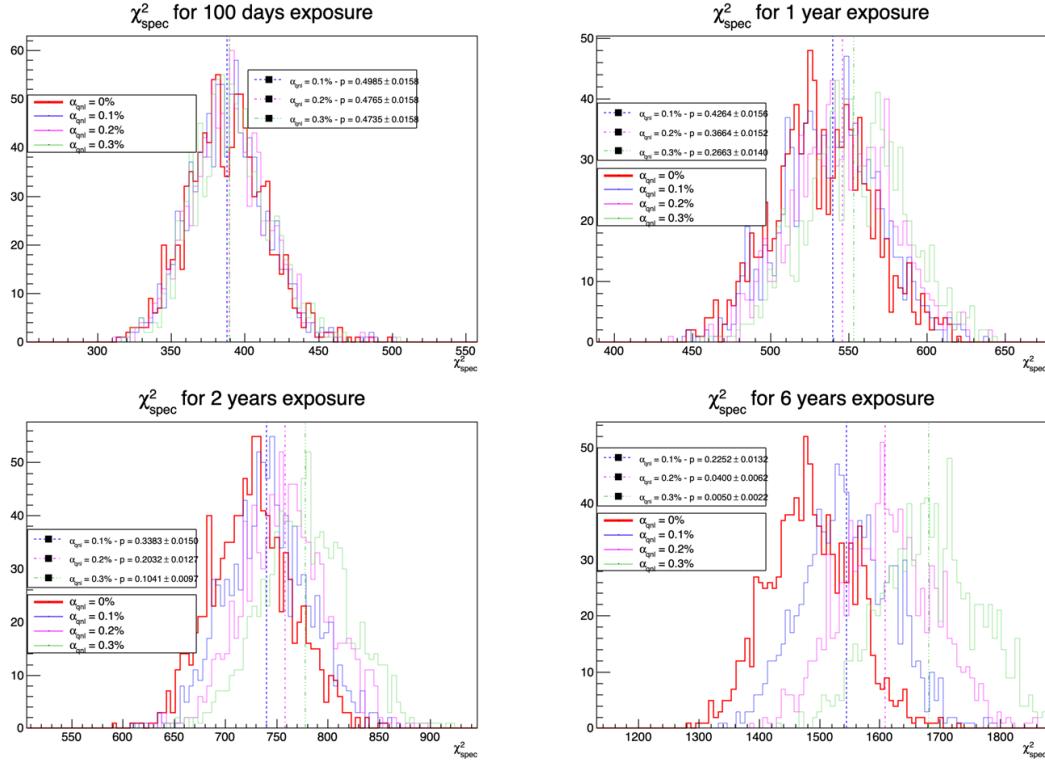


FIGURE 7.18 – Distribution of the  $\chi^2_{\text{spe}}$  for 1000 toys for different exposure. The dashed line represent the median of the distribution and the p-value are the percentage of the  $\alpha_{qnl} = 0$  distribution that are greater than those medians.

3947 This test relies solely on the estimated covariance matrix between the two spectra, requiring no  
 3948 fitting. As a result, it is a very lightweight test that can still provide valuable indications of potential  
 3949 unknown distortions between the two spectra.

3950 **Joint fit:**  $\chi^2_{H_0} - \chi^2_{H_1}$

3951 This test is the most complex, requiring two fit and the covariance matrix between the LPMT and  
 3952 SPMT spectra. The results are presented in Figure 7.19.

3953 The results are good, close to the  $\chi^2_{\text{spe}}$ , one with a p-value at 6 years for  $\alpha_{qnl} = 0.3\%$  of  $0.01 \pm 0.003$ .  
 3954 This sensitivity is consistent with that of  $\chi^2_{\text{spe}}$ .

3955 **Comparison of the parameters  $\delta \sin^2(2\theta_{12})$  and  $\delta \Delta m_{21}^2$**

3956 We can see that the  $\delta \Delta m_{21}^2$  has a very small discriminative power (Figure 7.21) even at 6 years  
 3957 exposure with a p-value of  $0.34 \pm 0.01$  for  $\alpha_{qnl} = 0.3\%$ . On the other hand  $\delta \theta_{12}$  (Figure 7.20) has  
 3958 much more discriminative power with a p-value for  $\alpha_{qnl} = 0.3\%$  of  $0.025 \pm 0.005$ . This test with  
 3959 a single joint fit seems to be still less powerful than the  $\chi^2_{\text{spe}}$ . This can be explained as this method  
 3960 only get information through the oscillation parameters  $\theta_{12}$  and  $\Delta m_{21}^2$  missing potential informations  
 3961 contained in  $\Delta m_{31}^2$ .

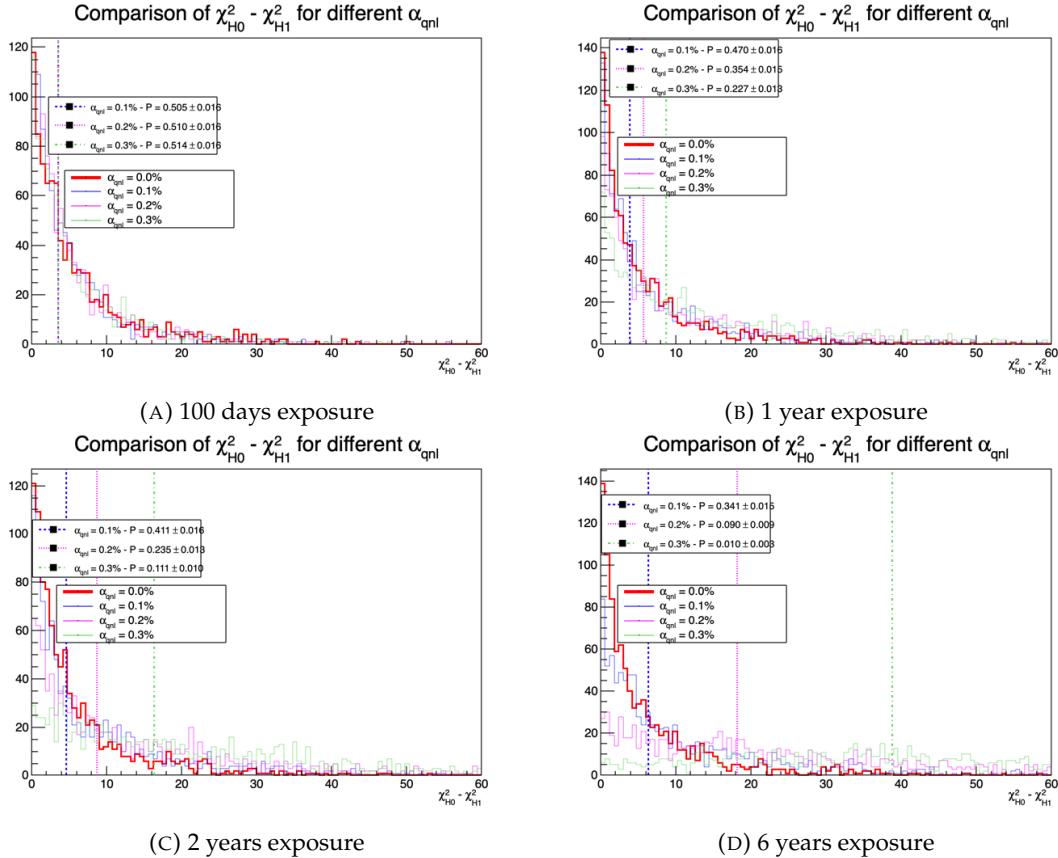


FIGURE 7.19 – Distribution of  $\chi^2_{H_0} - \chi^2_{H_1}$  for 1000 toys for different exposure. The dashed line represent the median of the distribution and the p-value are the percentage of the  $\alpha_{qnl} = 0$  distribution that are greater than those medians.

## 3962 Summary

The p-values from the different test and comparison for  $\alpha_{qnl} = 0.3\%$  are reported in Table 7.6.

	100 days	1 year	2 years	6 years
$\chi^2_{fid}$	0.49	0.41	0.29	0.090
$\chi^2_{spec}$	0.47	0.27	<b>0.10</b>	<b>0.005</b>
$\chi^2_{H_0} - \chi^2_{H_1}$	0.51	0.23	<b>0.11</b>	<b>0.010</b>
Comparison of $\delta \sin^2(2\theta_{12})$	0.39	0.2	0.14	0.025

TABLE 7.6 – Report of the p-value of the different tests and comparisons for  $\alpha_{qnl} = 0.3\%$  for the different exposures.

3963

## 3964 7.8 Conclusion and perspectives

3965 In this chapter, we present the development of a fit framework that allows us to fit multiple spectra  
3966 simultaneously. We also introduce a set of tools that enable us to detect potential distortions in one of

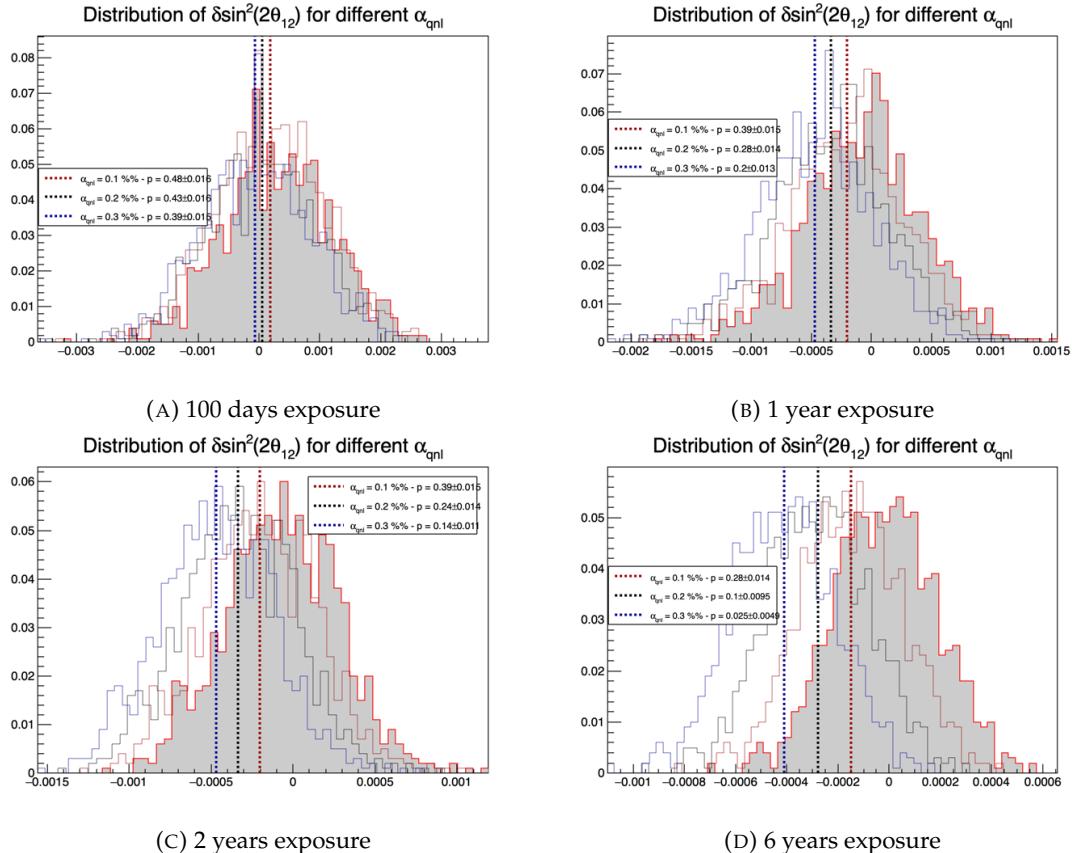


FIGURE 7.20 – Distribution of the  $\delta \sin^2(2\theta_{12})$  for 1000 toys for different exposure. The dashed line represent the median of the distribution and the p-value are the percentage of the  $\alpha_{qnl} = 0$  distribution that are greater than those medians.

3967 the two spectra. As an illustration of the capability of these tools, we use supplementary event-wise  
 3968 non-linearity and compare it to the potential residual event-wise non-linearity after calibration.

3969 Table 7.6 gives a synthetic view of the strength of our methods. As expected, two methods that  
 3970 exploit the knowledge of the correlations between the SPMT and LPMT spectra obtain the best  
 3971 results. At high exposures, if the QNL effects are not calibrated out as well as expected ( $> 0.3\%$ ),  
 3972 our best test statistics will be likely to detect them (median p-values below 10% after 2 years of data  
 3973 taking, about 1% after 6 years). In case of major effect (QNL or another unexpected instrumental  
 3974 effect) is worse, the detection will be even more likely. Below two years of data taking, only large  
 3975 unexpected instrumental effects can be detected.

3976 One of JUNO most important goals is to determine the NMO independently of other experiments.  
 3977 This should not happen before 6 years of data taking. Our results demonstrate that dual calorimetry  
 3978 with neutrino oscillation can be a useful approach to help ensure the robustness of this result.

### 3979 7.8.1 Empirical correlation matrix from fully simulated event

3980 As already explained several times, one of the limitation of this work is that we assume the SPMT and  
 3981 LPMT energy reconstructions to be totally uncorrelated. In reality, this is not true. The  $V$  covariance  
 3982 matrix used in the test statistics should therefore be evaluated accounting for this. This involves  
 3983 complications that make the subject out of the scope of this thesis. We present here a brief study

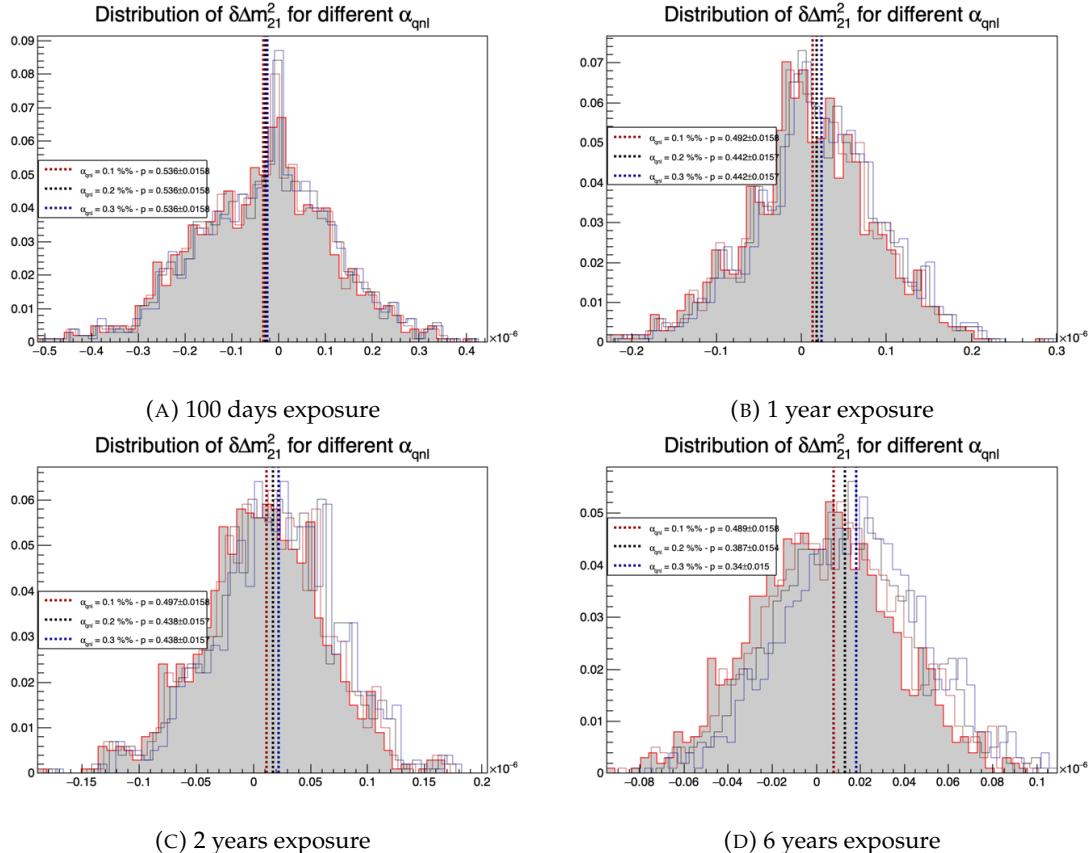


FIGURE 7.21 – Distribution of the  $\delta\Delta m_{21}^2$  for 1000 toys for different exposure. The dashed line represent the median of the distribution and the p-value are the percentage of the  $\alpha_{qnl} = 0$  distribution that are greater than those medians.

which goal is to get a rough idea of the impact of these reconstruction correlations.

The core of the idea is that the LPMT and SPMT reconstruction errors is bound to be correlated due to systematic effects. The first and most obvious one, for example, is energy escaping from the central detector. If the positron, or one of the two annihilation gamma, escape from the detector, less energy is deposited thus both of the systems will reconstruct a lower energy that was actually deposited. On a more subtle scale, the randomness in the production of scintillation photons is common for the two systems, if the liquid scintillator produces fewer scintillation photons for an event, both systems are likely to underestimate the energy.

We study those effects by computing from a dataset of IBD events, uniformly distributed in the CD, the correlation between the reconstruction errors on the energy

$$\text{Corr}(E_{rec}^{lpmt} - E_{vis}, E_{rec}^{spmt} - E_{vis}) \quad (7.32)$$

where  $E_{rec}^{lpmt}$  and  $E_{rec}^{spmt}$  are the reconstructed energies from both systems and  $E_{vis}$  the true visible energy. The OMILREC algorithm, presented in section 3.3, is used for the LPMT reconstruction  $E_{rec}^{lpmt}$ , and the CNN presented in Chapter 4 for the SPMT reconstruction  $E_{rec}^{spmt}$ .

The results of those correlations are presented in Figure 7.22 for the single energy and the interaction radius dependency, and Figure 7.23 for the dual energy and interaction radius dependencies.

The first observation here is that in most of the detector volume, the correlation between the SPMT

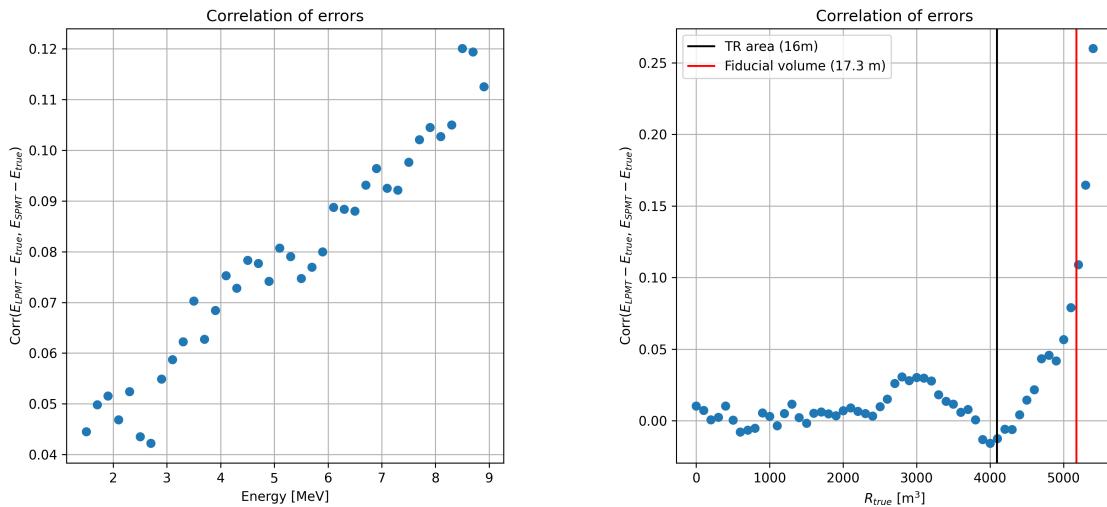


FIGURE 7.22 – Correlation on the reconstruction error between the LPMT and SPMT system as a function of (On the left) the energy, (On the right) the radius. The SPMT reconstruction comes from the NN presented in Chapter 4 and the LPMT reconstruction comes from OMILREC presented in Section 3.3. To prevent effect due to the CNN bad reconstruction, we select the event with  $1 < E_{dep} < 9$  MeV.

and LPMT energy reconstructions does not exceed a few percents, and is in general positive.

In principle, this correlation must be dominated by the fluctuations of the photon yield produced in the scintillator, which dominates the stochastic term of the resolution (see Equation 7.19). Indeed, in a given event, both the LPMT and SPMT reconstruct the energy from the same photon yield and both are affected in the same way by a fluctuation. The correlation is reduced by the fact that SPMT system, due to its low coverage, detect only a very small fraction of the photon. This sampling is also a random phenomenon : the corresponding fluctuations hide to some extent the fluctuations of the original photon yield, and are essentially independent of the random number of photons sampled by the LPMT.

When energy is deposited at high R, close to or in the total reflection area, the proximity of the PMTs increases the number of photons detected by LPMT, and therefore reduce the sampling fluctuations. In this case, the fluctuation of the original photon yield is less shuffled by the sampling fluctuations and the resulting correlation between the LPMT and SPMT reconstruction reaches high values, up to 25% (Fig. 7.22, right).

The original photon yield grows with the visible energy. For the same reason as above, the correlation grows as well, albeit far more slowly than as a function of  $R^3$ . On Fig. 7.23, one can see that cumulating the effects of high energy and high R, correlations can reach 35%. However, in the fiducial volume and at energies below 7 MeV (ie in a part of the spectrum containing the sensitivity to  $\Delta m_{12}^2$  and  $\sin^2(2\theta_{12})$ ), it never exceeds 15%.

To re-evaluate V with these reconstruction correlations accounted for, we should perform an empiric evaluation (like in Section 7.5.2). It would be based on toys generated with the IBD generator (see point 9 of Section 7.3.3), replacing the two independent random gaussian drawings by a drawing according to a 2 dimensional gaussian describing the  $(E_{rec}^{lpmt} - E_{vis})$  vs.  $(E_{rec}^{spmt} - E_{vis})$  distribution, and involving the correlations studied above.

A way must be found to include the variation of the correlation as a function of R and the  $E_{vis}$ . We have tried to define 2-dimensional regions in these variables, and defined each time the correspond-

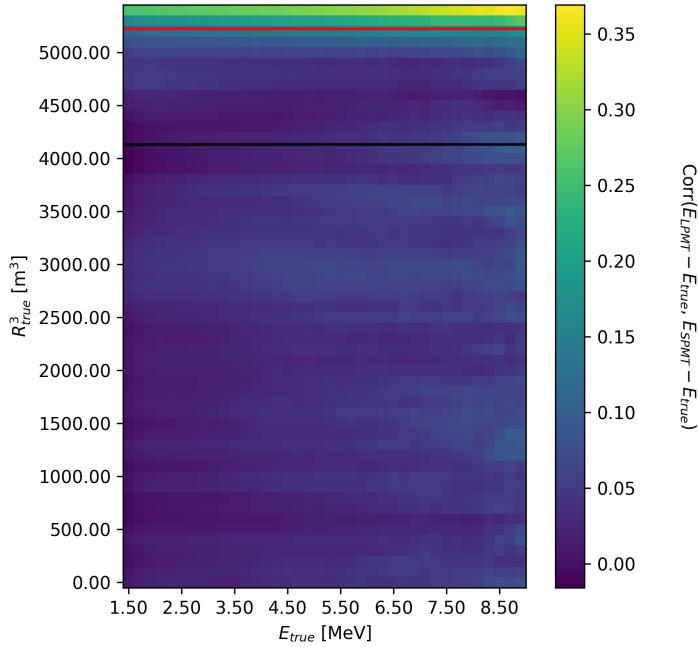


FIGURE 7.23 – Correlation on the reconstruction error between the LPMT and SPMT system as a function of the energy and the radius. The SPMT reconstruction comes from the NN presented in Chapter 4 and the LPMT reconstruction comes from OMILREC presented in Section 3.3. To prevent effect due to the CNN bad reconstruction, we select the event with  $1 < E_{dep} < 9 \text{ MeV}$ .

ing 2 dimensional gaussian. Then, we tuned the IBD generator to choose which of these gaussians to sample, based on the generated values of  $E_{vis}$  and  $R$ . Unfortunately, due to the limited statistics of the full simulation sample, the  $E_{vis} : R^3$  regions were too wide. It lead to sawtooth variations of the correlation and mean values of the gaussian between neighbouring regions. The reconstructed spectra finally showed irregularities instead of a normal, smooth aspect, making it improper for any oscillation analysis.

Before a solution can be found to this problem, we limit our conclusions to this :

- The correlation between the LPMT and SPMT energy reconstruction is positive. Therefore, the SPMT and LPMT spectra should be more correlated than assumed in the statistical tests presented in this chapter. With a proper treatment, we can therefore expect a higher sensitivity to unexpected instrumental effects like QNL.
- In 80% of the detector's volume, reconstruction correlations are low, and should not impact much the sensitivities of our test statistics. If the dependence of the correlation on  $E_{vis}$  and  $R$  proved too difficult to model, one could cut IBDs reconstructed in the Total reflection area. The loss in statistics would be limited, as well as the impact on the sensitivities of the statistical tests.

Additionally, this study is preliminary, as the background was neglected in the distortion test, and no systematic uncertainties were considered. Those points could be easily addressed by regenerating background spectra using the same reference as used by JUNO for the common inputs and by regenerating the systematic covariances matrix with both LPMT and SPMT spectra.

- 4046 The supplementary non-linearity was introduced event-wise but should be applied channel-wise to  
4047 account for the detector's non-uniformity. This can be addressed via generating oscillated spectra  
4048 through the JUNO official simulation. This process is very time consuming and require technical  
4049 development but could be achievable given enough time.
- 4050 The correlation matrix between the LPMT and SPMT spectra should also be further analyzed, as  
4051 indicated by the discrepancies between the theoretical and empirical correlation matrices.



# 4052 Conclusion

4053 The field of neutrino physics still has a lot of unanswered questions, namely the mass of the mass  
 4054 states, the Neutrino Mass Ordering (NMO), the possible existence of CP violation in the lepton sector,  
 4055 the unitarity of the PMNS oscillation matrix, and even the nature of the neutrino-Dirac or Majorana-  
 4056 is still unknown. To answer all of these questions, neutrino physics must advance into an era of  
 4057 precision measurements, of which JUNO will be a part.

4058 This thesis presents my contributions to the JUNO experiment. Its main goals are the measurement  
 4059 of the oscillation parameters  $\theta_{12}$ ,  $\Delta m_{21}^2$ , and  $\Delta m_{31}^2$  at the permille level, and to determine the Neutrino  
 4060 Mass Ordering with a significance that requires to reconstruct the energy of the reactor antineutrinos  
 4061 with a very high precision, and to understand this reconstruction very well. All my contributions  
 4062 are related to these goals.

4063

4064 In the first two chapters, I gave a short introduction to Neutrino physics and presented the JUNO  
 4065 experiment. I presented both the detector and various fit approaches used at JUNO to perform the  
 4066 reactor antineutrino oscillation analysis. It's a base to understand the fit I developed in Chapter 7.

4067 A large part of my thesis work was devoted to the development of Machine Learning algorithms  
 4068 for the reconstruction of reactor antineutrinos. In Chapter 3, I gave an introduction to a few types  
 4069 of algorithms (CNN, GNN) used at JUNO and in this thesis. I also present the existing antineutrino  
 4070 reconstruction methods, with and without machine learning, which are an important point of com-  
 4071 parison with the methods I developed during this thesis. I showed that the performance of the ML  
 4072 algorithms developed before the beginning of this thesis did not exceed in a convincing way the  
 4073 performance of JUNO's canonical likelihood based reconstruction algorithms.

4074

4075 In Chapter 4, I present the first algorithm I developed. It's a CNN reconstructing antineutrinos using  
 4076 only the SPMT system. Providing an alternative to classical methods in this context is interesting in  
 4077 its own right.

4078 It was also for me a gallop of test to learn about JUNO's environment. Finally, classical algorithms  
 4079 not being available in JUNO's public software, I could use this CNN in Chapter 7, where the SPMT  
 4080 reconstruction was necessary. The performance reached by this tool is close to that of classical  
 4081 methods as far as the energy is concerned, but worse when it comes to the reconstruction of the  
 4082 interaction position.

4083 One of the difficulties of my algorithm is that it has to train on a lot of pixels that have not been hit.  
 4084 This problem, partially due to the planar projection of a spherical experiment, is amplified by the  
 4085 specificities SPMTs (low coverage). The information these pixels carry is meaningless, which should  
 4086 cause problems in information aggregation. It could be solved by transforming the time information,  
 4087 a scalar, into a supplementary dimension in the image, resulting in the stacking of successive planar  
 4088 projections, each representing a time slice of the event. This would hopefully allow to match classical  
 4089 performances. I did not have enough time to implement such solutions, before I had to switch to  
 4090 my main thesis subjects. I also performed a combination of the CNN and the classical algorithm. Its  
 4091 performance exceeds that of the classical algorithm, demonstrating that there must exist an algorithm  
 4092 better using the input information.

4093 In Chapter 5, we formulated the hypothesis according to which ML or DL methods might yield  
 4094 better performance than the classical one if they manage to use more of the information present in  
 4095 the detector, by starting from a rawer level of data (PMT waveforms). Dealing with such a quantity  
 4096 of data requires architectures that help the network to identify essential information and to converge  
 4097 toward the result. We studied the potential of a GNN with an innovative architecture (heterogeneous  
 4098 Graph). It required a lot of technical developments, and a lot of work on the optimisation of the  
 4099 architecture and hyperparameters. This is the ML related work on which I provided most my efforts.  
 4100

4101 The best performance we obtained does not match that of the classical algorithm nor of other ML  
 4102 methods. We studied elements that suggest that when the GNN aggregates the signals from indi-  
 4103 vidual PMTs belonging to a certain region of the sphere, useful information, in particular temporal,  
 4104 is lost. This demonstrates the difficulty to find ML architectures that will actually improve recon-  
 4105 struction performance. Future versions of my GNN will have to work on this. We can look for new  
 4106 ways to link various regions of the detector, and spend further time refining and adapt the message  
 4107 passing algorithm.

4108  
 4109 In Chapter 6, we worked on ML reliability. We believe that the first step to ensure the reliability of  
 4110 the reconstruction is to benefit of a variety of algorithms. The combination method developed during  
 4111 this thesis allow to not only compare performance and behavior but also to probe in the difference  
 4112 in information used. This also underlines the interest of developing several algorithms for the same  
 4113 tasks, which are then useful even when they do not reach the best performance. However, this is  
 4114 possible only if all algorithms are available to any user. For that reason, my first work on reliability  
 4115 was to implement in JUNO's common software some tools necessary to include in the ML algorithms  
 4116 until then developed as standalone tools, available only to their authors. I also implemented one of  
 4117 these ML algorithms.

4118 We know it is crucial for JUNO not only to reconstruct very precisely the energy of antineutrinos,  
 4119 but also to understand the quality of this reconstruction, and the differences in this between real data  
 4120 and the models assumed by the fits employed to perform the oscillation analysis. We suspect that  
 4121 some subtle differences in the charge and time measured by individual PMTs could affect JUNO's  
 4122 results by distorting very slightly the energy spectrum, while being invisible to data/Monte Carlo  
 4123 comparisons carried out with calibration or signal free control samples. In this chapter 6, I also  
 4124 discuss the exploration of the usage of an Adversarial Neural Network which goal is to help identify  
 4125 the kind of discrepancies that could have this effect, by generating perturbations to the charge and  
 4126 time measured by individual PMTs.

4127 The conclusion of this part explains that this first ANN prototype does not manage to generate  
 4128 perturbations that affect IBD events more than control sample events. However, this exploration  
 4129 taught us several things, among which : it is very difficult to design an ANN able to introduce  
 4130 perturbations at the individual PMT level; some physics-informed guidance will be necessary to  
 4131 obtain an operational tool in the future.

4132  
 4133 The last chapter of this thesis is devoted to Dual calorimetry. There are several concrete applications  
 4134 of this technique. Generically, it is based on the comparison of quantities reconstructed individually  
 4135 by the LPMT and the SPMT systems. It will be used at calibration level. In this thesis, we explore  
 4136 another way, called Dual Calorimetry analysis with neutrino oscillation. It exploits the potential  
 4137 discrepancies between oscillation analyses carried out with either PMT systems.

4138 We designed four statistical tests to detect unexpected instrumental effects in one of the systems or  
 4139 both. We evaluated their sensitivity to a concrete problem: the Charge non linearity (QNL) that will  
 4140 plausibly affect LPMTs. These tests are : the direct comparison of the values of  $\sin^2(2\theta_{12})$  and  $\Delta m_{21}^2$   
 4141 obtained with the LPMT system or the SPMT system ; a direct comparison of the energy spectra

4142 reconstructed by either systems ; and two other tests based on a joint fit of these spectra. A crucial  
4143 ingredient there are the correlations between these spectra, which exist even at the level of statistical  
4144 uncertainties. We designed ways to evaluate them.

4145 We observe that the most powerful tests are those which indeed fully account for these correlations  
4146 : unexpected instrumental effects are not detected only because data spectra do not match the  
4147 predicted spectra but also because they are not consistent with the predicted correlations.

4148 JUNO's most important result will concern the determination of the NMO with JUNO's data only,  
4149 i.e. independent of other experiments. A 3 sigma result is possible with about 6 years of data taking.  
4150 With such statistics, our best statistical tests should detect with a p-value around 1% a QNL effect  
4151 if the calibration phase has not corrected it as well as expected. It proves the interest of the Dual  
4152 calorimetry analysis with neutrino oscillation.

4153 Several assumptions have been discussed concerning the impact of systematic uncertainties, of the  
4154 backgrounds or of the correlation between the SPMT and LPMT reconstructions. They will be the  
4155 subject of future works to make Dual Calorimetry with neutrino oscillation fully operational. We do  
4156 not expect the sensitivities observed here to change much after these refinements.

4157 This work was also the occasion of important technical developments which constitute a major  
4158 improvement of the analysis framework the Subatech group will use to contribute to JUNO's results.



<sup>4159</sup> **Appendix A**

<sup>4160</sup> **Calculation of optimal  $\alpha$  for estimator combination**

<sup>4162</sup> This annex the details of the determination of the optimal  $\alpha$  for estimator combination presented in  
<sup>4163</sup> section 4.3.2.

<sup>4164</sup> As a reminder, the combined estimator  $\hat{\theta}$  of  $X$  is defined as

$$\hat{\theta}(X) = \alpha\theta_N + (1 - \alpha)\theta_C; \alpha \in [0; 1] \quad (\text{A.1})$$

<sup>4165</sup> where  $\theta_N$  and  $\theta_C$  are both estimator of  $X$ .

<sup>4166</sup> **A.1 Unbiased estimator**

For the unbiased estimator, it is straight-forward. We search  $\alpha$  such as  $E[\hat{\theta}] = X$

$$E[\hat{\theta}] = E[\alpha\theta_N + (1 - \alpha)\theta_C] \quad (\text{A.2})$$

$$= E[\alpha\theta_N] + E[(1 - \alpha)\theta_C] \quad (\text{A.3})$$

$$= \alpha E[\theta_N] + (1 - \alpha)E[\theta_C] \quad (\text{A.4})$$

$$= \alpha(\mu_N + X) + (1 - \alpha)(\mu_C + X) \quad (\text{A.5})$$

$$X = \alpha\mu_N + \mu_C - \alpha\mu_C + X \quad (\text{A.6})$$

$$0 = \alpha(\mu_N - \mu_C) + \mu_C \quad (\text{A.7})$$

$$(A.8)$$

$$\Rightarrow \alpha = \frac{\mu_C}{\mu_C - \mu_N} \quad (\text{A.9})$$

<sup>4167</sup> **A.2 Optimal variance estimator**

The  $\alpha$  for this estimator is a bit more tricky. By expanding the variance we get

$$\text{Var}[\hat{\theta}] = \text{Var}[\alpha\theta_N + (1 - \alpha)\theta_C] \quad (\text{A.10})$$

$$= \text{Var}[\alpha\theta_N] + \text{Var}[(1 - \alpha)\theta_C] + \text{Cov}[\alpha(1 - \alpha)\theta_N\theta_C] \quad (\text{A.11})$$

$$= \alpha^2\sigma_N^2 + (1 - \alpha)^2\sigma_C^2 + 2\alpha(1 - \alpha)\sigma_N\sigma_C\rho_{NC} \quad (\text{A.12})$$

<sup>4168</sup> where, as a reminder,  $\rho_{NC}$  is the correlation factor between  $\theta_C$  and  $\theta_N$ .

Now we try to find the minima of  $\text{Var}[\hat{\theta}]$  with respect to  $\alpha$ . For this we evaluate the derivative

$$\frac{d}{d\alpha} \text{Var}[\hat{\theta}] = 2\alpha\sigma_N^2 - 2(1-\alpha)\sigma_C^2 + 2\sigma_N\sigma_C\rho_{NC}(1-2\alpha) \quad (\text{A.13})$$

$$= 2\alpha(\sigma_N^2 + \sigma_C^2 - 2\sigma_N\sigma_C\rho_{NC}) - 2\sigma_C^2 + 2\sigma_N\sigma_C\rho_{NC} \quad (\text{A.14})$$

then find the minima and maxima of this derivative by evaluating

$$\frac{d}{d\alpha} \text{Var}[\hat{\theta}] = 0 \quad (\text{A.15})$$

$$2\alpha(\sigma_N^2 + \sigma_C^2 - 2\sigma_N\sigma_C\rho_{NC}) - 2\sigma_C^2 + 2\sigma_N\sigma_C\rho_{NC} = 0 \quad (\text{A.16})$$

$$2\alpha(\sigma_N^2 + \sigma_C^2 - 2\sigma_N\sigma_C\rho_{NC}) = 2\sigma_C^2 - 2\sigma_N\sigma_C\rho_{NC} \quad (\text{A.17})$$

$$\alpha = \frac{\sigma_C^2 - \sigma_N\sigma_C\rho_{NC}}{\sigma_N^2 + \sigma_C^2 - 2\sigma_N\sigma_C\rho_{NC}} \quad (\text{A.18})$$

4169 This equation shows only one solution which is a minima. From Eq. A.18 arise two singularities:

4170 —  $\sigma_N = \sigma_C = 0$ . This is not a problem because as physicists we never measure with an absolute  
4171 precision, neither us or our detectors are perfect.

4172 —  $\sigma_N = \sigma_C$  and  $\rho_{CN} = 1$ . In this case  $\theta_C$  and  $\theta_N$  are the same estimator in term of variance thus  
4173 any value for  $\alpha$  yield the same result: an estimator with the same varianve as the original ones.

<sup>4174</sup> **Appendix B**

## <sup>4175</sup> Charge spherical harmonics analysis

<sup>4176</sup> When looking at JUNO events we can clearly see some pattern in the charge repartition based on  
<sup>4177</sup> the event radius as illustrated in figure B.4. When dealing with identifying features and pattern on a  
<sup>4178</sup> spherical plane, the astrophysics community have been using, with success, the spherical harmonic  
<sup>4179</sup> decomposition. The principle is similar to a frequency analysis via Fourier transform. It comes to  
<sup>4180</sup> saying that a function  $f(r, \theta, \phi)$ , here our charge repartition of the spherical plane constructed by our  
<sup>4181</sup> PMTs, can be expressed

$$f(r, \theta, \phi) = \sum_{l=0}^{\infty} \sum_{m=-l}^l a_l^m r^l Y_l^m(\theta, \phi) \quad (\text{B.1})$$

<sup>4182</sup> where  $a_l^m$  are constants complex factor,  $Y_l^m(\theta, \phi) = Ne^{im\phi} P_l^m(\cos \theta)$  are the spherical harmonics of  
<sup>4183</sup> degree  $l$  and order  $m$  and  $P_l^m$  their associated Legendre Polynomials. Those harmonics are illustrated  
<sup>4184</sup> in figure B.1. By reducing the problem to the unit sphere  $r = 1$ , we get rid of the term  $r^l$ . The Healpix  
<sup>4185</sup> library [127] offer function to efficiently find the  $a_l^m$  factor from a given Healpix map.

<sup>4186</sup> For the above decomposition, we will define the *Power* of an harmonic as

$$S_{ff}(l) = \frac{1}{2l+1} \sum_{m=-l}^l |a_l^m|^2 \quad (\text{B.2})$$

<sup>4187</sup> and the *Relative Power* as:

$$P_l^h = \frac{S_{ff}(l)}{\sum_l S_{ff}(l)} \quad (\text{B.3})$$

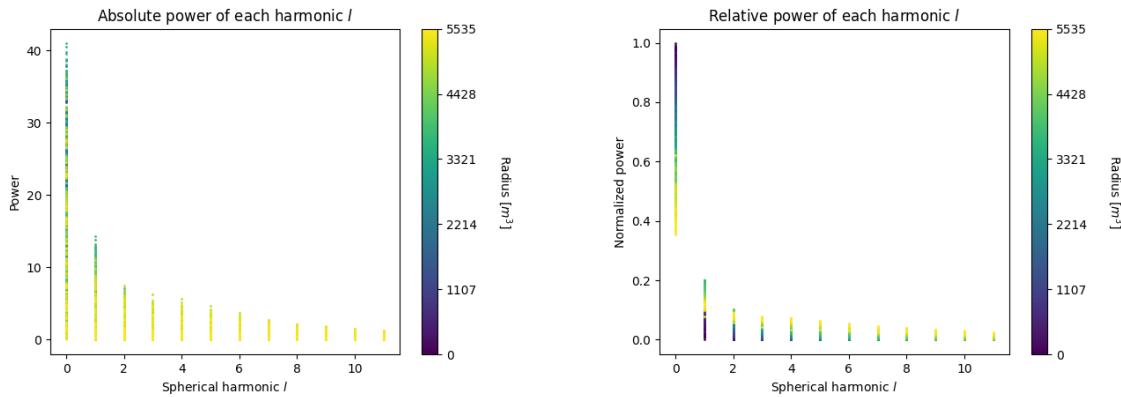
<sup>4188</sup> For this study we will use 10k positron events with  $E_{kin} \in [0; 9]$  MeV uniformly distributed in the  
<sup>4189</sup> CD from the JUNO official simulation version J23.0.1-rc8.dc1 (released the 7th January 2024). All the  
<sup>4190</sup> event are *calib* level, with simulation of the physics, electronics, digitizations and triggers. We first  
<sup>4191</sup> take a sub-set of 1k events and look at the power and relative power distribution depending on the  
<sup>4192</sup> radius and harmonic degree  $l$ . The results are shown in figure B.2. While don't see any pattern in  
<sup>4193</sup> absolute power, it is pretty clear that there is a correlation between the relative power of  $l = 0$  and  
<sup>4194</sup> the radius of the event.

<sup>4195</sup> When applying the same study but dependent on the energy, no clear correlation appear. The results  
<sup>4196</sup> for the  $l = 0$  harmonic are presented in the figure B.5. Thus, in this study we will focus on the radial  
<sup>4197</sup> dependency of the relative power of each harmonic.

<sup>4198</sup> In figures B.6 and B.7 are presented the distribution of the relative power of each harmonic for  $l \in$   
<sup>4199</sup>  $[0, 11]$ . The relation between the radius and the relative power become even more clear, especially  
<sup>4200</sup> for the first harmonics  $l \in [0, 4]$ . After that for  $l > 4$  their relative power is close to 0 for central event,  
<sup>4201</sup> thus loosing power. It also interesting to note the change of behavior in the TR area, clearly visible  
<sup>4202</sup> for  $l = 1$  and  $l = 2$ .

$l:$	$P_\ell^m(\cos \theta) \cos(m\varphi)$	$P_\ell^{ m }(\cos \theta) \sin( m \varphi)$	
0 s			
1 p			
2 d			
3 f			
4 g			
5 h			
6 i			
$m:$	6 5 4 3 2 1 0	-1 -2 -3 -4 -5 -6	

FIGURE B.1 – Illustration of the real part of the spherical harmonics

FIGURE B.2 – Scatter plot of the absolute and relative power, respectively on the left and right plot, of each harmonic degree  $l$ . The color indicate the radius of the event.

As an erzats of reconstruction algorithm, we fit each of those distribution with a 9th degree polynomial which give us the relation

$$F(R^3) \longmapsto P_l^h \quad (\text{B.4})$$

We do it this way because some of the distribution have multiple solution for a given relative power, for example  $l = 1$ , while each radius give only one power. We now just need to find

$$F^{-1}(P_l^h) \longmapsto R^3 \quad (\text{B.5})$$

Inverting a 9th degree polynomial is hard, if not impossible. The presence of multiple roots for the same power complexify the task even more. To circumvent this problem, we reconstruct the radius by locating the minima of  $(F(R^3) - \hat{P}_l^h)^2$  where  $\hat{P}_l^h$  is the measured power fraction.

To distinguish between multiple possible minima, we use as a starting point the radius given by the procedure on  $l = 0$  that, by looking at the fit in figure B.6, should only present one minima. For  $l > 0$  we also impose bound on the possible reconstructed  $R^3$  as  $R^3 \in [R_0^3 - 100, R_0^3 + 100]$  where  $R_0^3$  is the reconstructed  $R^3$  by the harmonic  $l = 0$ .

4214 The minimization algorithm used are the Bent algorithm for  $l = 0$  and the Bounded algorithm for  
 4215  $l > 0$  provided by the Scipy library [143]. We then do the mean of the reconstructed radius from  
 4216 the different harmonics. The reconstruction results are shown in figure B.3. The performance seems  
 4217 correct but we see heavy fluctuation in the bias. To really be used as a reconstruction algorithm, the  
 4218 method needs to be refined as discussed in the next section.

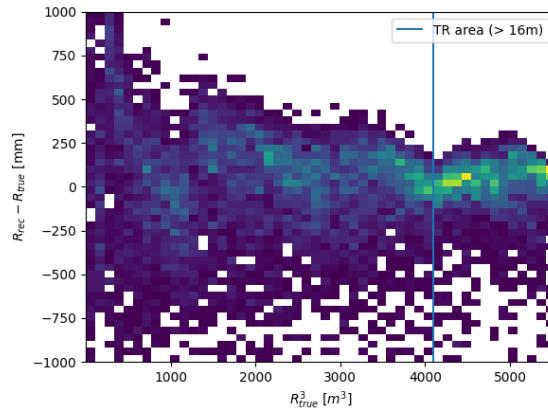


FIGURE B.3 – Error on the reconstructed radius vs the true radius by the harmonic method

## Conclusion

4219 We have clearly shown in this analysis the relevance the of relative harmonic power for radius  
 4220 reconstruction, and provided an erzats of a reconstruction algorithm. We will not delve further  
 4221 in this thesis but if we wanted to refine this algorithm multiple paths can be explored:

- 4222 — No energy signature in the harmonics: This is surprising that there is no correlation between  
 4223 the energy and the amplitude of the harmonics. We know that the energy is heavily correlated  
 4224 with the total number of photoelectrons collected, it would be unintuitive that we see no  
 4225 relation.
- 4226 — Localization of the event: We shown here the relation between the relative power of the har-  
 4227 monic and the radius but don't get any information about the  $\theta$  and  $\phi$  spherical coordinates.  
 4228 This information is probably hidden in the individual power of each order  $m$  of the degree  $l$ .  
 4229 This intuition comes from the figure B.1 where in the higher degree  $l$  we see that the order  $m$   
 4230 are oriented. Intuitively, the order should be able to indicate a direction where the signal is  
 4231 more powerful.
- 4232 — Combination of the degree power: Here we combined the radius reconstructed by the different  
 4233 degree via a simple mean but we shown in section 4.3.2 and annex A that this is note the optimal  
 4234 way to combine estimator. A more refined algorithm probably exist to take into account the  
 4235 predicting power of each order.

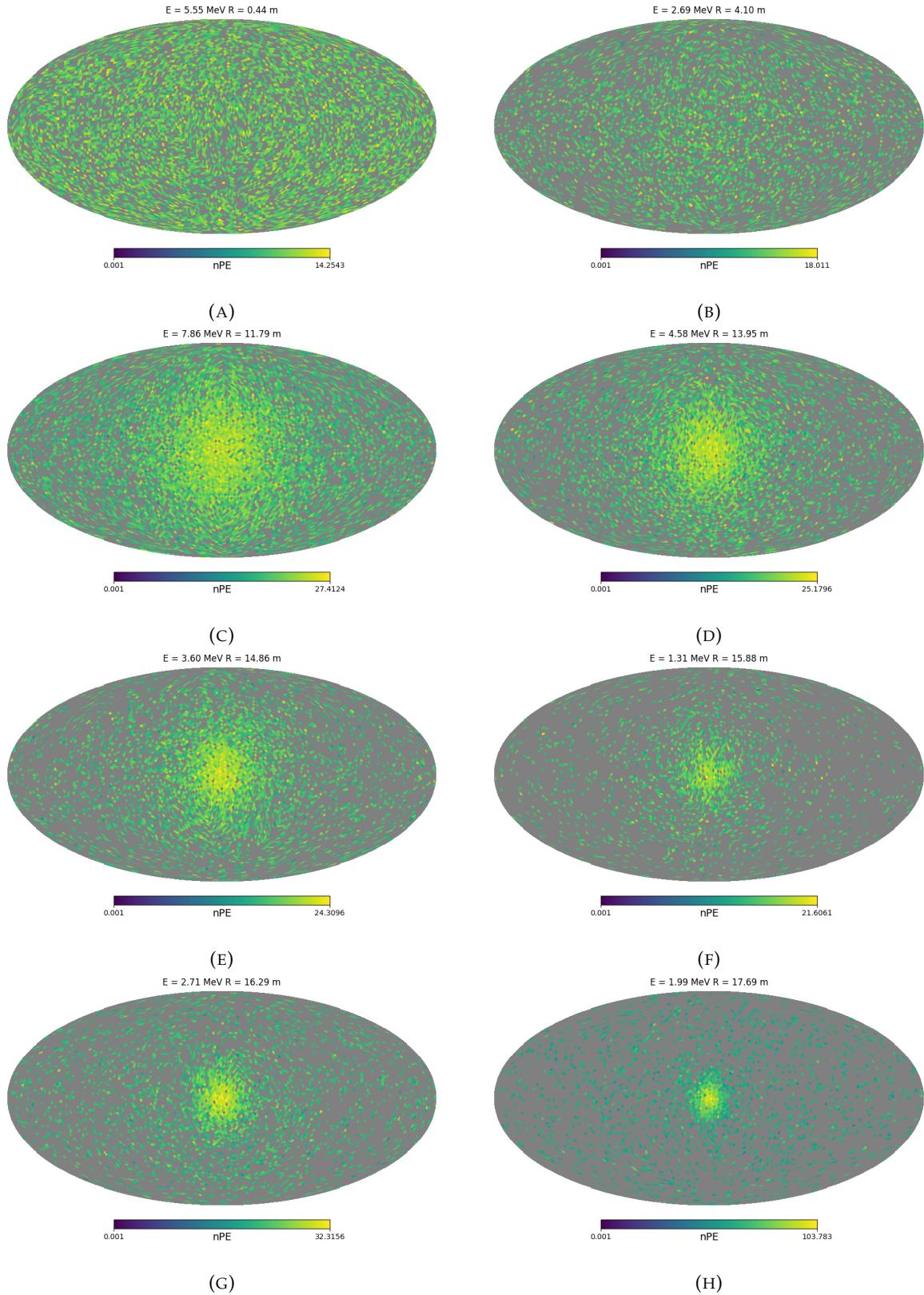


FIGURE B.4 – Charge repartition in JUNO as seen by the Healpix segmentation. Those are Healpix map of order 5 (i.e. 12288 pixels). The color represent the summed charge of the PMTs in each pixels. The color scale is logarithmic. The view have been centered to prevent event deformations.

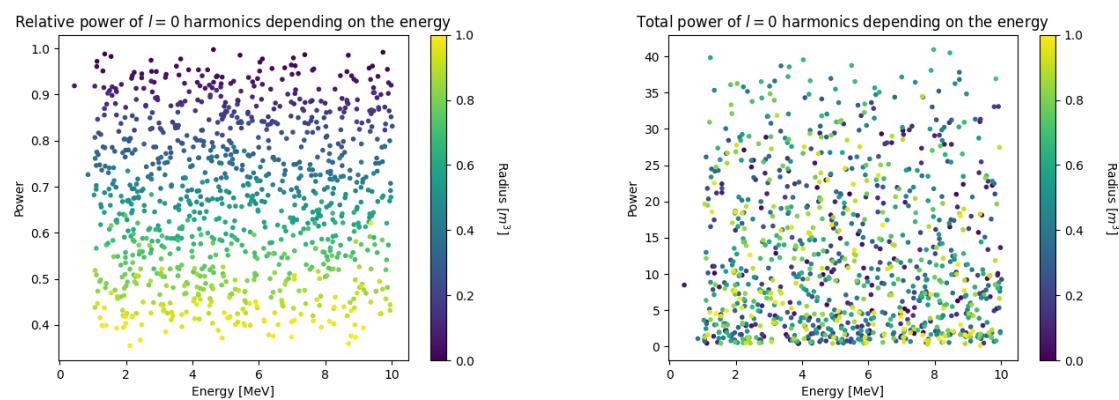


FIGURE B.5 – Scatter plot of the absolute and relative power, respectively on the left and right plot, of the  $l = 0$  harmonic. The color indicate the radius of the event.

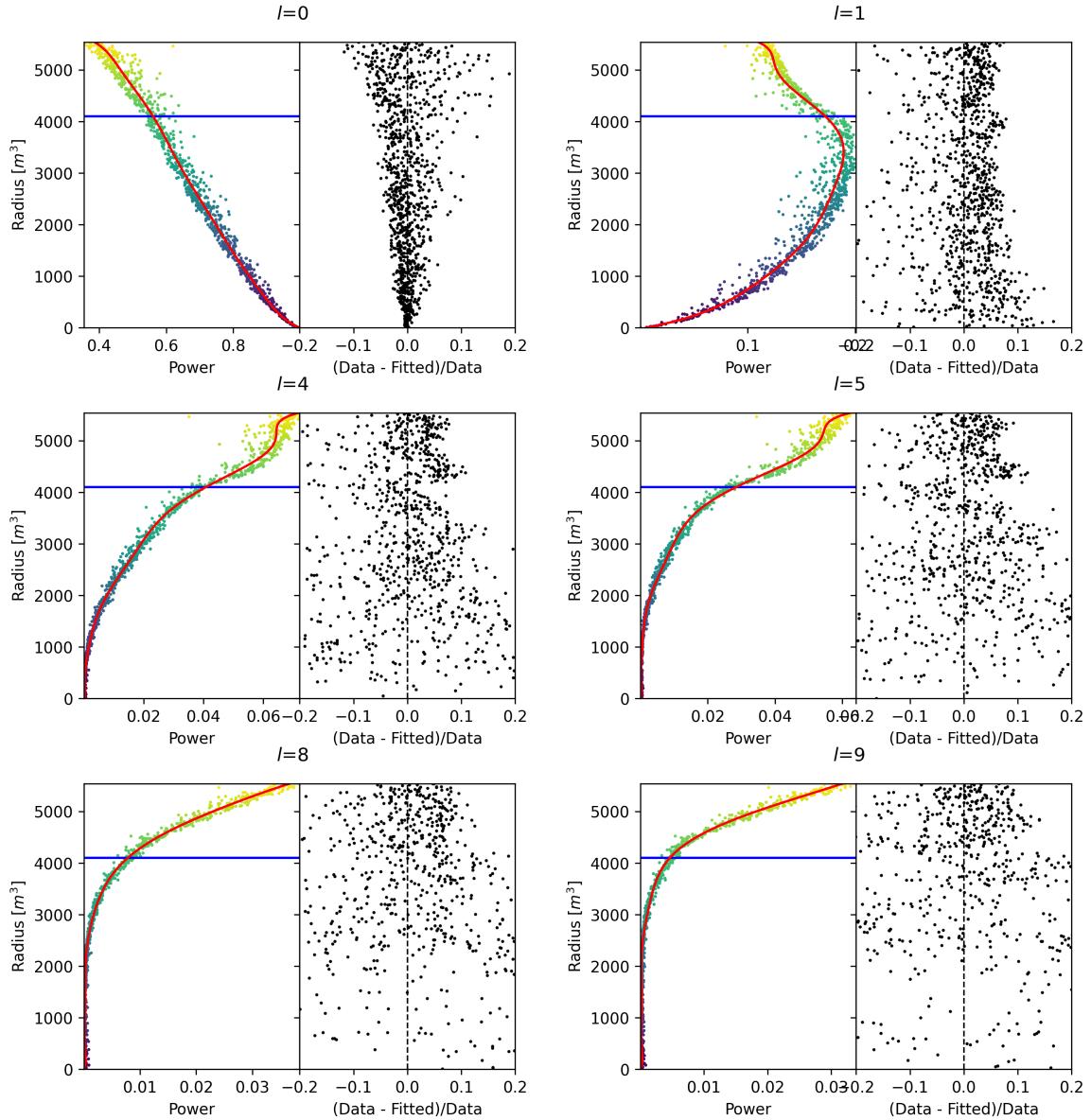


FIGURE B.6 – Plot of the distribution of the relative power of each harmonic dependent on  $R^3$  (on the left). The Total Reflection (TR) area is represented by the horizontal blue line. The distribution are fitted using a 9th degree polynomial (red curve). The relative power error between the distribution and the fit is represented on the left. **Part 1**

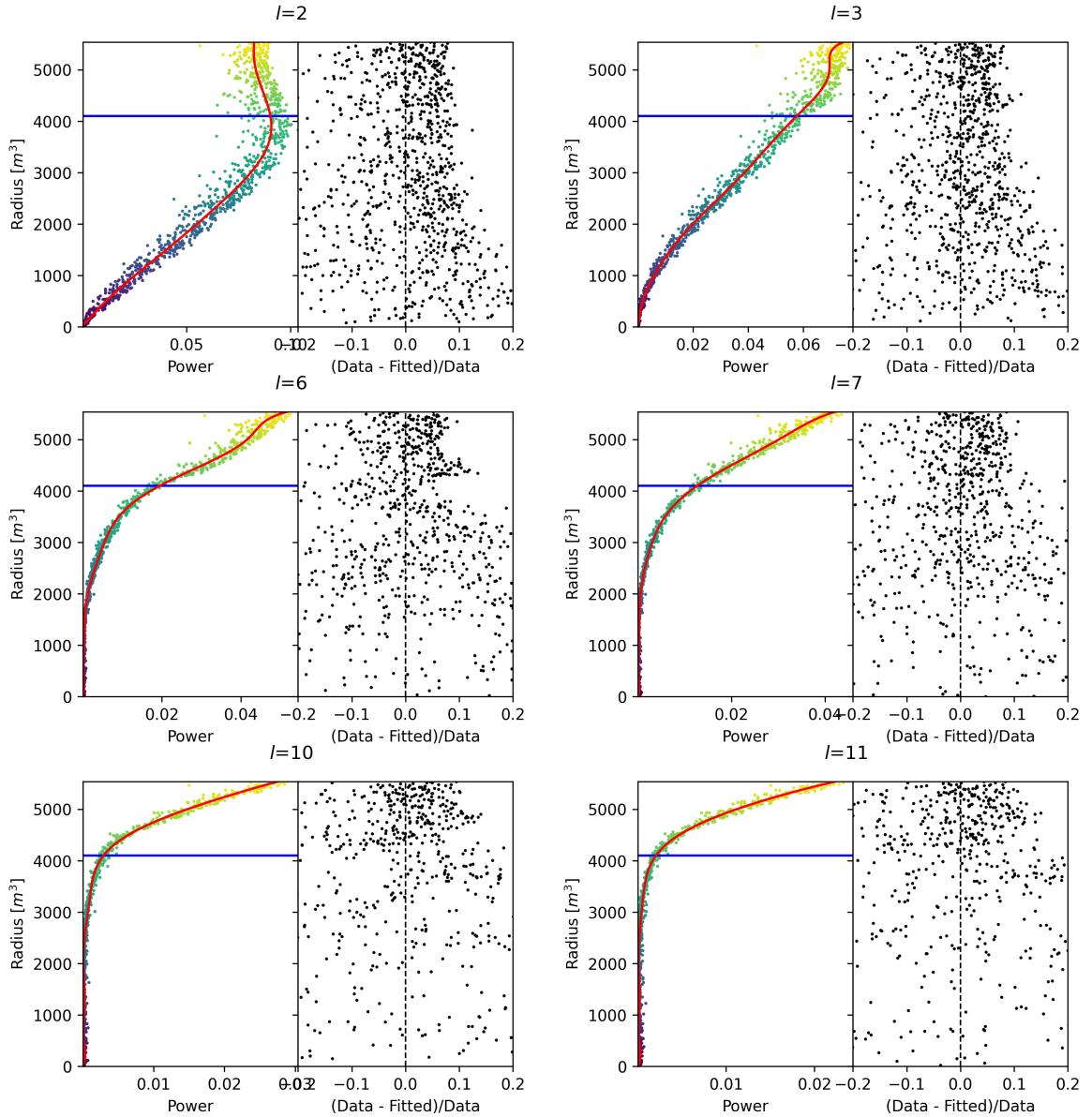


FIGURE B.7 – Plot of the distribution of the relative power of each harmonic dependent on  $R^3$  (on the left). The Total Reflection (TR) area is represented by the horizontal blue line. The distribution are fitted using a 9th degree polynomial (red curve). The relative power error between the distribution and the fit is represented on the left. **Part 2**



<sup>4237</sup> **Appendix C**

<sup>4238</sup> **Correction of  $E_{vis}$  bias**

<sup>4239</sup> The reconstruction algorithms that are presented in this thesis in Chapters 4 and 5 do not reconstruct  
<sup>4240</sup> the same energy as the classical algorithms presented in section 3.3. Our algorithms reconstruct the  
<sup>4241</sup> deposited energy  $E_{dep}$  while the classical algorithms reconstruct a visible energy  $E_{vis}$ .

To understand this phenomena, let's look at the equation 3.27:

$$\hat{\mu}(r, \theta, \theta_{pmt}, E_{vis}) = \frac{1}{E_{vis}} \frac{1}{M} \sum_i^M \frac{\frac{\bar{Q}_i}{\bar{Q}_i} - \mu_i^D}{DE_i}, \quad \mu_i^D = DNR_i \cdot L$$

<sup>4242</sup> which define the expected  $N_{pe}/E$ . This define a linear relation between the number of photoelectrons  
<sup>4243</sup> and the energy. However we discussed in sections 2.3.2 and 2.4 that the number of photoelectrons  
<sup>4244</sup> collected by the LPMT system do not follow a linear relationship. Thus this visible energy is not  
<sup>4245</sup> linear with the deposited energy. This effect is corrected in physics analysis and in Chapter 7 by  
<sup>4246</sup> applying the calibrated non-linearity profile the energy spectrum.

<sup>4247</sup> When we need to compare our algorithm that reconstruct the deposited energy to the classical  
<sup>4248</sup> algorithms we need to correct this non-linearity. For this we fit the systematic bias of the classical  
<sup>4249</sup> algorithm using a 5th degree polynomial

$$\frac{E_{dep}}{E_{vis}} = \sum_{i=0}^5 P_i E_{dep}^i \quad (C.1)$$

<sup>4250</sup> The fitted distribution and the corresponding fit is presented in figure C.1. The value fitted for this  
<sup>4251</sup> correction are presented in table C.1.

$P_0$	$1.24541 +/- 0.00585121$
$P_1$	$-0.168079 +/- 0.00716387$
$P_2$	$0.0489947 +/- 0.00312875$
$P_3$	$-0.00747111 +/- 0.000622003$
$P_4$	$0.000570998 +/- 5.7296e-05$
$P_5$	$-1.72588e-05 +/- 1.98355e-06$

TABLE C.1 – Parameters of the 5th degree polynomial used to correct Omilrec reconstructed energy.

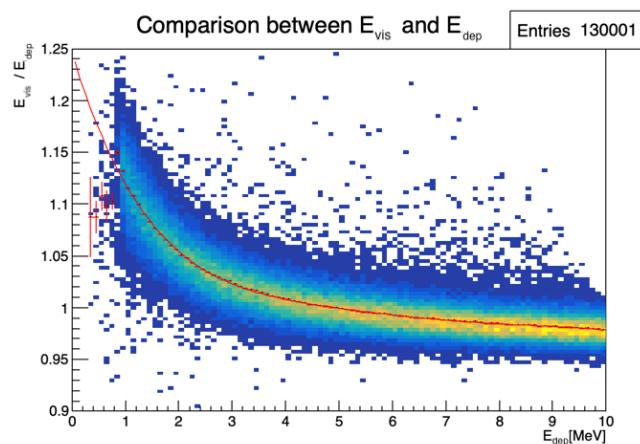


FIGURE C.1 – Comparison between Omilrec reconstructed  $E_{vis}$  and the deposited energy  $E_{dep}$ . The profile of the distribution  $E_{vis}/E_{dep}$  vs  $E_{dep}$  is fitted with a 5th degree polynomial.

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# List of Abbreviations

<b>ACU</b>	Automatic Calibration Unit
<b>ANN</b>	Adversarial Neural Network
<b>BDT</b>	Boosted Decision Tree
<b>BFP</b>	Best Fit Point
<b>CD</b>	Central Detector
<b>CLS</b>	Cable Loop System
<b>CNN</b>	Convolutional NN
<b>DNN</b>	Deep NN
<b>DN</b>	Dark Noise
<b>EDM</b>	Event Data Model
<b>FCDNN</b>	Fully Connected Deep NN
<b>GNN</b>	Graph NN
<b>GT</b>	Guiding Tube
<b>IBD</b>	Inverse Beta Decay
<b>IO</b>	Inverse Ordering
<b>JUNO</b>	Jiangmen Underground Neutrino Observatory
<b>LPMT</b>	Large PMT
<b>LR</b>	Learning Rate
<b>LS</b>	Liquid Scintillator
<b>MC</b>	Monte Carlo simulation
<b>ML</b>	Machine Learning
<b>MSE</b>	Mean Squared Error
<b>NMO</b>	Neutrino Mass Ordering
<b>NN</b>	Neural Network
<b>NO</b>	Normal Ordering
<b>NPE</b>	Number of Photo Electron
<b>OSIRIS</b>	Online Scintillator Internal Radioactivity Investigation System
<b>PE</b>	Photo Electron
<b>PMT</b>	Photo-Multipliers Tubes
<b>PRelu</b>	Parametrized Rectified Linear Unit
<b>QNL</b>	Charge (Q) Non Linearity
<b>ROV</b>	Remotely Operated under-LS Vehicle
<b>ReLU</b>	Rectified Linear Unit
<b>ResNet</b>	Residual Network
<b>SGD</b>	Stochastic Gradient Descent
<b>SPMT</b>	Small PMT
<b>TAO</b>	Taishan Antineutrino Observatory
<b>TR Area</b>	Total Reflexion Area
<b>TTS</b>	Time Transit Spread
<b>TT</b>	Top Tracker
<b>UWB</b>	Under Water Boxes
<b>WCD</b>	Water Cherenkov Detector



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**Titre :** Méthode Deep Learning and analyse Double Calorimétrique pour la mesure de haute précision des paramètres d'oscillation des neutrinos dans JUNO

**Mot clés :** Neutrinos; expérience JUNO; Deep Learning; reconstruction d'IBD; oscillations des neutrinos; double calorimetrie

**Résumé :** JUNO est un observatoire de neutrinos à scintillateur liquide, polyvalent et medium baseline (environ 52 km), situé en Chine. Ses principaux objectifs sont de mesurer les paramètres d'oscillation  $\theta_{12}$ ,  $\Delta m_{21}^2$  et  $\Delta m_{31}^2$  avec une précision au pour-mille et de déterminer l'ordre des masses des neutrinos avec un niveau de confiance de  $3\sigma$ . Atteindre ces objectifs nécessite une résolution énergétique sans précédent de  $3\%/\sqrt{E(\text{MeV})}$  avec cette technologie. Cela demande une compréhension approfondie des divers effets au sein du détecteur. Le

système de double calorimetrie, composé de deux systèmes de mesure distincts observant le même événement, permet non seulement une calibration mais aussi une détection des effets du détecteur avec une grande précision, comme démontré dans cette thèse. Le Deep Learning, un outil de plus en plus utilisé en physique expérimentale, joue un rôle crucial dans cet effort. Dans cette thèse, je présente le développement, l'application et l'analyse des techniques de Deep Learning pour la reconstruction d'évènements dans l'expérience JUNO.

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**Title:** Deep learning methods and Dual Calorimetric analysis for high precision neutrino oscillation measurements at JUNO

**Keywords:** Neutrinos; JUNO experiment; Deep learning; IBD reconstruction; neutrinos Oscillation; dual Calorimetry

**Abstract:** JUNO is a multipurpose, medium baseline ( $\sim 52$  km) liquid scintillator neutrino observatory located in China. Its primary objectives are to measure the oscillation parameters  $\theta_{12}$ ,  $\Delta m_{21}^2$ , and  $\Delta m_{31}^2$  with per mil precision and to determine the neutrino mass ordering at a  $3\sigma$  confidence level. Achieving these goals requires an unprecedented energy resolution of  $3\%/\sqrt{E(\text{MeV})}$  with this technology. This demands a comprehensive understanding of the various effects within the

detector. The Dual Calorimetry system-two distinct measurement systems observing the same event-enables not only high-precision calibration but also detection of detector effects, as demonstrated in this thesis. Deep learning, an increasingly powerful tool in physics, plays a critical role in this effort. In this thesis, I present the development, application, and analysis of Deep Learning techniques for reconstruction in the JUNO experiment.

