

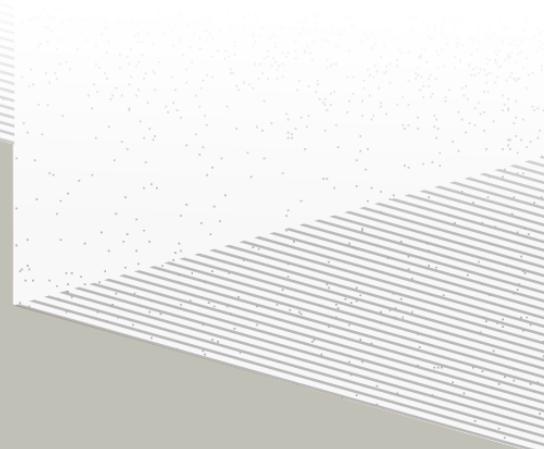
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# THÈSE DE DOCTORAT DE

L'UNIVERSITÉ DE NANTES

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*Matière, Molécules, Matériaux*  
Spécialité : *Physique des particules*



Par

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**Precision measurement of solar neutrino oscillation parameters  
with the JUNO small PMTs system and test of the unitarity of the  
PMNS matrix**

Thèse présentée et soutenue à Nantes, le Too soon and too early at the same time  
Unité de recherche : Laboratoire SUBATECH, UMR 6457

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|              |                                       |
|--------------|---------------------------------------|
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## <sup>3</sup> List of Abbreviations

|                |   |
|----------------|---|
| <b>ACU</b>     | Automatic Calibration Unit                |
| <b>CD</b>      | Central Detector                          |
| <b>CLS</b>     | Cable Loop System                         |
| <b>GT</b>      | Guiding Tube                              |
| <b>IBD</b>     | Inverse Beta Decay                        |
| <b>IO</b>      | Inverse Ordering                          |
| <b>JUNO</b>    | Jiangmen Underground Neutrino Observatory |
| <b>LPMT</b>    | Large PMT                                 |
| <b>LS</b>      | Liquid Scintillator                       |
| <b>MC</b>      | Monte Carlo simulation                    |
| <b>ML</b>      | Machine Learning                          |
| <b>NMO</b>     | Neutrino Mass Ordering                    |
| <b>NN</b>      | Neural Network                            |
| <b>NO</b>      | Normal Ordering                           |
| <b>PE</b>      | Photo Electron                            |
| <b>PMT</b>     | Photo-Multipliers Tubes                   |
| <b>ROV</b>     | Remotely Operated under-LS Vehicle        |
| <b>SPMT</b>    | Small PMT                                 |
| <b>TR Area</b> | Total Reflexion Area                      |
| <b>TTS</b>     | Time Transit Spread                       |
| <b>TT</b>      | Top Tracker                               |
| <b>WCD</b>     | Water Cherenkov Detector                  |



# 4 Contents

|    |   |    |
|----|---|----|
| 5  | <b>Contents</b>   | 1  |
| 6  | <b>Remerciements</b>  | 3  |
| 7  | <b>Introduction</b>   | 5  |
| 8  | <b>1 Neutrino physics</b>   | 7  |
| 9  | 1.1 Standard model . . . . .  | 7  |
| 10 | 1.1.1 Limits of the standard model . . . . .                                  | 7  |
| 11 | 1.2 Historic of the neutrino . . . . .  | 7  |
| 12 | 1.3 Oscillation . . . . .   | 7  |
| 13 | 1.3.1 Phenomologies . . . . .   | 7  |
| 14 | 1.4 Open questions . . . . .  | 7  |
| 15 | <b>2 The JUNO experiment</b>  | 9  |
| 16 | 2.1 Neutrinos physics in JUNO . . . . .                                       | 9  |
| 17 | 2.1.1 Reactor neutrino oscillation for NMO and precise measurements . . . . . | 10 |
| 18 | 2.1.2 Other physics . . . . .   | 12 |
| 19 | 2.2 The JUNO detector . . . . .   | 14 |
| 20 | 2.2.1 Principle of detection . . . . .  | 14 |
| 21 | 2.2.2 Central Detector (CD) . . . . .   | 15 |
| 22 | 2.2.3 Veto detector . . . . .   | 18 |
| 23 | 2.3 Calibration strategy . . . . .  | 18 |
| 24 | 2.3.1 Energy scale calibration . . . . .                                      | 19 |
| 25 | 2.3.2 Calibration system . . . . .  | 19 |
| 26 | 2.4 Software . . . . .  | 20 |
| 27 | 2.5 State of the art of the Offline IBD reconstruction in JUNO . . . . .      | 21 |
| 28 | 2.5.1 Interaction vertex reconstruction . . . . .                             | 22 |
| 29 | 2.5.2 Energy reconstruction . . . . .   | 26 |
| 30 | 2.5.3 Machine learning for reconstruction . . . . .                           | 28 |
| 31 | 2.6 JUNO sensitivity to NMO and precise measurements . . . . .                | 29 |
| 32 | 2.6.1 Fitting procedure . . . . .   | 29 |
| 33 | <b>3 Machine learning and Artificial Neural Network</b>                       | 31 |
| 34 | 3.1 History of the Machine learning . . . . .                                 | 31 |
| 35 | 3.2 Boosted Decision Tree (BDT) . . . . .                                     | 31 |

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|    |   |    |
|----|---|----|
| 36 | <b>3.3 Artificial Neural Network (NN) . . . . .</b>                               | 31 |
| 37 | 3.3.1 Fully Connected Deep Neural Network (FCDNN) . . . . .                       | 31 |
| 38 | 3.3.2 Convolutional Neural Network (CNN) . . . . .                                | 31 |
| 39 | 3.3.3 Graph Neural Network (GNN) . . . . .  | 31 |
| 40 | 3.3.4 Adversarial Neural Network (ANN) . . . . .                                  | 31 |
| 41 | <b>4 Image recognition for IBD reconstruction with the SPMT system</b>            | 33 |
| 42 | <b>5 Graph representation of JUNO for IBD reconstruction with the LPMT system</b> | 35 |
| 43 | <b>6 Reliability of machine learning methods</b>                                  | 37 |
| 44 | <b>7 Discrimination of e+/e- events in JUNO</b>                                   | 39 |
| 45 | <b>8 Conclusion</b>   | 41 |
| 46 | <b>Bibliography</b>   | 43 |

<sup>47</sup> **Remerciements**



# 48 Introduction



<sup>49</sup>

## Chapter 1

<sup>50</sup>

# Neutrino physics

<sup>51</sup>

*The neutrino, or  $\nu$  for the close friends, a fascinating and invisible particle. Some will say that dark matter also have those property but at least we are pretty confident that neutrinos exists.*

<sup>52</sup>

## 1.1 Standard model

<sup>53</sup>

### 1.1.1 Limits of the standard model

<sup>54</sup>

## 1.2 Historic of the neutrino

<sup>55</sup>

### First theories

<sup>56</sup>

### Discovery

<sup>57</sup>

### Milestones and anomalies

<sup>58</sup>

## 1.3 Oscillation

<sup>59</sup>

### 1.3.1 Phenomologies

<sup>60</sup>

## 1.4 Open questions

Decrire le m  
Regarder th  
Kochebina  
Limite du r  
Interessant,  
les neutrino  
CP ? Pb des



<sup>61</sup> **Chapter 2**

<sup>62</sup> **The JUNO experiment**

<sup>63</sup> “Ave Juno, rosae rosam, et spiritus rex”. It means nothing but I found it in tone.

<sup>64</sup> The first idea of a medium baseline ( $\sim 52$  km) experiment, was explored in 2008 [1] where it was  
<sup>65</sup> demonstrated that the Neutrino Mass Ordering (NMO) could be determined by a medium baseline  
<sup>66</sup> experiment if  $\sin^2(2\theta_{13}) > 0.005$  without the requirements of accurate knowledge of the reactor  
<sup>67</sup> antineutrino spectra and the value of  $\Delta m_{32}^2$ . From this idea is born the Jiangmen Underground  
<sup>68</sup> Neutrino Observatory (JUNO) experiment.

<sup>69</sup> JUNO is a neutrino detection experiment under construction located in China. Its main objectives  
<sup>70</sup> are the determination of the mass ordering at the  $3\text{-}4\sigma$  level in 6 years of data taking and the mea-  
<sup>71</sup> surement at the sub-percent precision of the oscillation parameters  $\Delta m_{21}^2$ ,  $\sin^2 \theta_{12}$ ,  $\Delta m_{32}^2$  and with less  
<sup>72</sup> precision  $\sin^2 \theta_{13}$ [2].



FIGURE 2.1 – On the left: Location of the JUNO experiment and its reactor sources in southern china. On the right: Aerial view of the experimental site

<sup>73</sup> For this JUNO will measure the electronic anti-neutrinos ( $\bar{\nu}_e$ ) flux coming from the nuclear reac-  
<sup>74</sup> tors of Taishan, Yangjiang, for a total power of  $26.6 \text{ GW}_{th}$ , and the Daya Bay power plant to a lesser  
<sup>75</sup> extent. Details about the power plants and there expected flux of  $\bar{\nu}_e$  can be found in the table 2.1. The  
<sup>76</sup> distance of 53 km has been specifically chosen to maximize the disappearance probability of the  $\bar{\nu}_e$ .

<sup>77</sup> **2.1 Neutrinos physics in JUNO**

<sup>78</sup> Even if JUNO design (section 2.2) was optimized for the measurement of the NMO, its large  
<sup>79</sup> detection volume, excellent energy resolution and background level and understanding make it also  
<sup>80</sup> an excellent detector to measure the flux coming from other neutrino sources. Thus the scientific  
<sup>81</sup> program of JUNO extends way over reactor antineutrinos. The following section is an overview of  
<sup>82</sup> the different physics topic JUNO will contribute in the coming years.

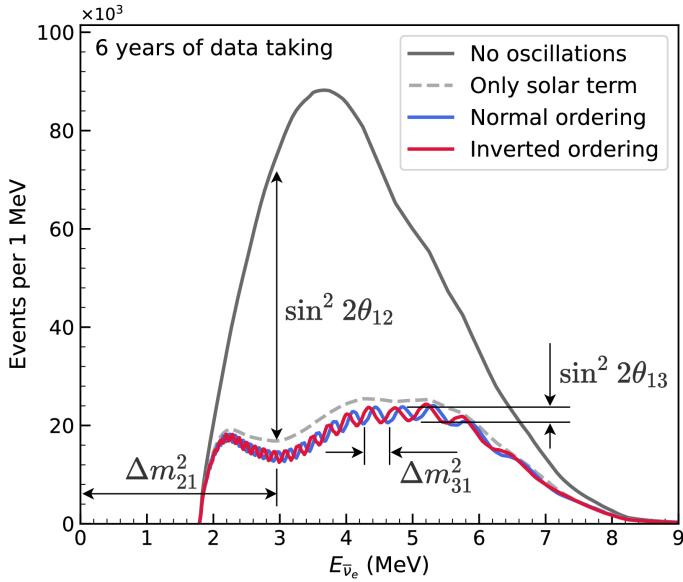


FIGURE 2.2 – Expected number of neutrinos event per MeV in JUNO after 6 years of data taking. The black curve shows the flux if there was no oscillation. The light gray curve shows the oscillation if only the solar terms are taken in account ( $\theta_{12}$ ,  $\Delta m_{21}^2$ ). The blue and red curve shows the spectrum in the case of, respectively, NO and IO. The dependency of the oscillation to the different parameters are schematized by the double sided arrows. We can see the NMO sensitivity by looking at the fine phase shift between the red and the blue curve.

### 83 2.1.1 Reactor neutrino oscillation for NMO and precise measurements

Previous works [1, 3] shows that oscillation parameters and the NMO can be observed by looking at the  $\bar{\nu}_e$  disappearance spectrum coming from medium baseline nuclear reactor. This disappearance probability can be expressed as [2] :

$$P(\bar{\nu}_e \rightarrow \bar{\nu}_e) = 1 - \sin^2 2\theta_{12} c_{13}^4 \sin^2 \frac{\Delta m_{21}^2 L}{4E} - \sin^2 2\theta_{13} \left[ c_{12}^2 \sin^2 \frac{\Delta m_{31}^2 L}{4E} + s_{12}^2 \sin^2 \frac{\Delta m_{32}^2 L}{4E} \right]$$

84 Where  $s_{ij} = \sin \theta_{ij}$ ,  $c_{ij} = \cos \theta_{ij}$ ,  $E$  is the  $\bar{\nu}_e$  energy and  $L$  is the baseline. We can see the sensitivity  
 85 to the NMO in the dependency to  $\Delta m_{32}^2$  and  $\Delta m_{31}^2$  causing a phase shift of the spectrum as we can  
 86 see in the figure 2.2. By carefully fitting this spectrum, one can extract the NMO and the oscillation  
 87 parameters. The fit is reviewed in more details in the section 2.6.1. To reach the desired sensitivity,  
 88 JUNO must meet multiple requirements but most notably:

- 89 1. An energy resolution of  $3\%/\sqrt{E(\text{MeV})}$  to be able to distinguish the fine structure of the fast  
 90 oscillation.
- 91 2. An energy precision of 1% in order to not err on the location of the oscillation pattern.
- 92 3. A baseline of  $53 \pm 0.5$  km to maximise the  $\bar{\nu}_e$  oscillation probability.
- 93 4. At least  $\approx 100,000$  events to limit the spectrum distortion due to statistical uncertainties.

### 94 $\bar{\nu}_e$ flux coming from nuclear power plants

95 To get such high measurements precision, it is necessary to have a very good understanding of  
 96 the sources characteristics. For its NMO and precise measurement studies, JUNO will observe the

97 energy spectrum of neutrinos coming from the nuclear power plants Taishan and Yangjiang's cores,  
 98 located at 53 km of the detector to maximise the disappearance probability of the  $\bar{\nu}_e$ .

| Reactor   | Power (GW <sub>th</sub> ) | Baseline (km) | IBD Rate (day <sup>-1</sup> ) | Relative Flux (%) |
|-----------|---------------------------|---------------|-------------------------------|-------------------|
| Taishan   | 9.2                       | 52.71         | 15.1                          | 32.1              |
| Core 1    | 4.6                       | 52.77         | 7.5                           | 16.0              |
| Core 2    | 4.6                       | 52.64         | 7.6                           | 16.1              |
| Yangjiang | 17.4                      | 52.46         | 29.0                          | 61.5              |
| Core 1    | 2.9                       | 52.74         | 4.8                           | 10.1              |
| Core 2    | 2.9                       | 52.82         | 4.7                           | 10.1              |
| Core 3    | 2.9                       | 52.41         | 4.8                           | 10.3              |
| Core 4    | 2.9                       | 52.49         | 4.8                           | 10.2              |
| Core 5    | 2.9                       | 52.11         | 4.9                           | 10.4              |
| Core 6    | 2.9                       | 52.19         | 4.9                           | 10.4              |
| Daya Bay  | 17.4                      | 215           | 3.0                           | 6.4               |

TABLE 2.1 – Characteristics of the nuclear power plants observed by JUNO. The IBD rate are estimated from the baselines, the reactors full thermal power, selection efficiency and the current knowledge of the oscillation parameters

99 The  $\bar{\nu}_e$  coming from reactors are emitted from  $\beta$ -decay of unstable fission fragments. The Taishan  
 100 and Yangjiang reactors are pressurised water reactor (PWR), the same type as Daya Bay. In those  
 101 type of reactor more the 99.7 % and  $\bar{\nu}_e$  are produced by the fissions of four fuel isotopes <sup>235</sup>U, <sup>238</sup>U,  
 102 <sup>239</sup>Pu and <sup>241</sup>Pu. The neutrino flux per fission of each isotope is determined by the inversion of the  
 103 measured  $\beta$  spectra of fission product [4–8] or by calculation using the nuclear databases [9, 10]. The  
 104 neutrino flux coming from a reactor at a time  $t$  can be predicted using

$$\phi(E_\nu, t)_r = \frac{W_{th}(t)}{\sum_i f_i(t) e_i} \sum_i f_i(t) S_i(E_\nu) \quad (2.1)$$

105 where  $W_{th}(t)$  is the thermal power of the reactor,  $f_i(t)$  is the fraction fission of the  $i$ th isotope,  $e_i$  its  
 106 thermal energy released in each fission and  $S_i(E_\nu)$  the neutrino flux per fission for this isotope. Using  
 107 this method, the flux uncertainty is expected to be of an order of 2-3 % [11].

108 In addition to those prediction, a satellite experiment named TAO[12] will be setup near the  
 109 reactor core Taishan-1 to measure with an energy resolution of 2% at 1 MeV the neutrino flux coming  
 110 from the core. It will help identifying unknown fine structure and give more insight on the  $\bar{\nu}_e$  flux  
 111 coming from this reactor.

112 One the open issue about reactor anti-neutrinos flux is the so-called neutrino anomaly [13], an  
 113 unexpected surplus of neutrino emission in the spectra around 5 MeV. Multiples scientists are trying  
 114 to explain this surplus by advanced recalculation of the nuclei model during beta decay [14, 15] but  
 115 no consensus on this issue has been reached yet.

## 116 Background in the neutrinos reactor spectrum

117 Considering the close reactor neutrinos flux as the main signal, the signals that are considered as  
 118 background are:

- 119 — The geoneutrinos producing background in the 0.511 ~ 2.7 MeV region.
- 120 — The neutrinos coming from the other nuclear reactors around Earth.

121 In addition to all those physics signal, non-neutrinos signal that would mimic an IBD will also be  
 122 present. It is composed of:

- 123 — The signal coming from radioactive decay ( $\alpha$ ,  $\gamma$ ,  $\beta$ ) from natural radioactive isotopes in the  
 124 material of the detector.
- 125 — Cosmogenic event such as fast neutrons and activated isotopes induced by muons passing  
 126 through the detector, most notably the spallation on <sup>12</sup>C.

<sup>127</sup> All those events represent a non-negligable part of the spectrum as shown in figure 2.3.

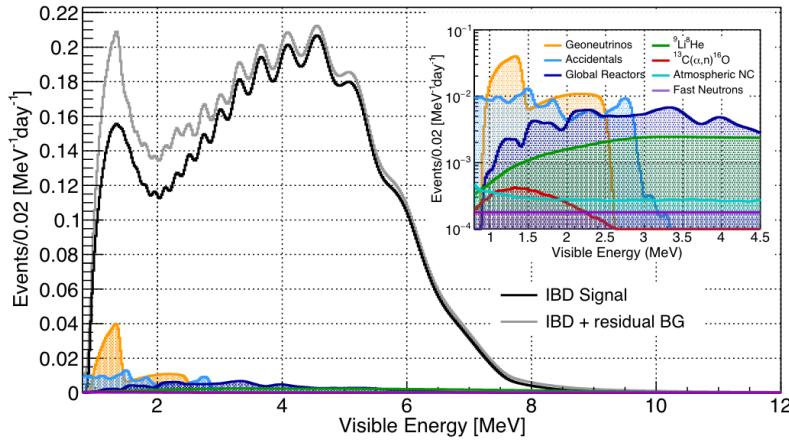


FIGURE 2.3 – Expected visible energy spectrum measured with the LPMT system with (grey) and without (black) backgrounds. The background amount for about 7% of the IBD candidate and are mostly localized below 3 MeV [11]

### <sup>128</sup> Identification of the mass ordering

<sup>129</sup> To identify the mass ordering, we fit the neutrino energy spectrum under the two hypothesis of  
<sup>130</sup> NO and IO. Those two fit give us two  $\chi^2$ , respectively  $\chi^2_{NO}$  and  $\chi^2_{IO}$ . By computing the difference  
<sup>131</sup>  $\Delta\chi^2 = \chi^2_{NO} - \chi^2_{IO}$  we can determine the most probable mass ordering: NO if  $\Delta\chi^2 > 0$  and IO  
<sup>132</sup> if  $\Delta\chi^2 < 0$ . Current studies shows that the expected sensitivity the mass ordering would be of  
<sup>133</sup>  $3.4\sigma$  after 6 years of data taking in nominal setup[2]. More detailed explanations about the fitting  
<sup>134</sup> procedure can be found in the section 2.6.1.

### <sup>135</sup> Precise measurement of the oscillations parameters

<sup>136</sup> The oscillations parameters  $\theta_{12}$ ,  $\theta_{13}$ ,  $\Delta m_{21}^2$ ,  $\Delta m_{31}^2$  are free parameters in the fit of the oscillation  
<sup>137</sup> spectrum. The precision on those parameters have been estimated and are shown in table 2.2. Wee  
<sup>138</sup> see that for  $\theta_{12}$ ,  $\Delta m_{21}^2$ ,  $\Delta m_{31}^2$ , precision at 6 years is better than the reference precision by an order of  
<sup>139</sup> magnitude [11]

|   | Central Value | PDG 2020            | 100 days            | 6 years              | 20 years            |
|---|---------------|---------------------|---------------------|----------------------|---------------------|
| $\Delta m_{31}^2 (\times 10^{-3} \text{ eV}^2)$ | 2.5283        | $\pm 0.034$ (1.3%)  | $\pm 0.021$ (0.8%)  | $\pm 0.0047$ (0.2%)  | $\pm 0.0029$ (0.1%) |
| $\Delta m_{21}^2 (\times 10^{-3} \text{ eV}^2)$ | 7.53          | $\pm 0.18$ (2.4%)   | $\pm 0.074$ (1.0%)  | $\pm 0.024$ (0.3%)   | $\pm 0.017$ (0.2%)  |
| $\sin^2 \theta_{12}$                            | 0.307         | $\pm 0.013$ (4.2%)  | $\pm 0.0058$ (1.9%) | $\pm 0.0016$ (0.5%)  | $\pm 0.0010$ (0.3%) |
| $\sin^2 \theta_{13}$                            | 0.0218        | $\pm 0.0007$ (3.2%) | $\pm 0.010$ (47.9%) | $\pm 0.0026$ (12.1%) | $\pm 0.0016$ (7.3%) |

TABLE 2.2 – A summary of precision levels fir the oscillation parameters. The reference value (PDG 2020 [16]) is compared with 100 days, 6 years and 20 years of JUNO data taking.

### <sup>140</sup> 2.1.2 Other physics

<sup>141</sup> While the design of JUNO is tailored to measure  $\bar{\nu}_e$  coming from nuclear reactor, JUNO will be  
<sup>142</sup> able to detect neutrinos coming from other sources thus allowing for a wide range of physics studies  
<sup>143</sup> as detailed in the table 2.3 and in the following sub-sections.

| Research             | Expected signal                      | Energy region | Major backgrounds          |
|----------------------|--------------------------------------|---------------|----------------------------|
| Reactor antineutrino | 60 IBDs/day                          | 0–12 MeV      | Radioactivity, cosmic muon |
| Supernova burst      | 5000 IBDs at 10 kpc                  | 0–80 MeV      | Negligible                 |
| DSNB (w/o PSD)       | 2300 elastic scattering              |               |                            |
| Solar neutrino       | 2–4 IBDs/year                        | 10–40 MeV     | Atmospheric $\nu$          |
| Atmospheric neutrino | hundreds per year for ${}^8\text{B}$ | 0–16 MeV      | Radioactivity              |
| Geoneutrino          | hundreds per year                    | 0.1–100 GeV   | Negligible                 |
|                      | $\approx 400$ per year               | 0–3 MeV       | Reactor $\nu$              |

TABLE 2.3 – Detectable neutrino signal in JUNO and the expected signal rates and major background sources

#### <sup>144</sup> **Geoneutrinos**

<sup>145</sup> Geoneutrinos designate the antineutrinos coming from the decay of long-lived radioactive elements inside the Earth. The 1.8 MeV threshold necessary for the IBD makes it possible to measure <sup>146</sup> geoneutrinos from  ${}^{238}\text{U}$  and  ${}^{232}\text{Th}$  decay chains. The studies of geoneutrinos can help refine the Earth <sup>147</sup> crust models but is also necessary to characterise their signal, as they are a background to the mass <sup>148</sup> ordering and oscillations parameters studies.

#### <sup>150</sup> **Atmospheric neutrinos**

<sup>151</sup> Atmospheric neutrinos are neutrinos originating from the decay of  $\pi$  and  $K$  particles that are <sup>152</sup> produced in extensive air showers initiated by the interactions of cosmic rays with the Earth atmosphere. Earth is mostly transparent to neutrinos below the PeV energy, thus JUNO will be able to <sup>153</sup> see neutrinos coming from all directions. Their baseline range is large (15km  $\sim$  13000km), they can <sup>154</sup> have energy between 0.1 GeV and 10 TeV and will contain all neutrino and antineutrinos flavour. <sup>155</sup> Their studies is complementary to the reactor antineutrinos and can help refine the constraints on <sup>156</sup> the NMO [2].

#### <sup>158</sup> **Supernovae burst neutrinos**

<sup>159</sup> Neutrinos are crucial component during all stages of stellar collapse and explosion. Detection <sup>160</sup> of neutrinos coming for core collapse supernovae will provide us important informations on the <sup>161</sup> mechanisms at play in those events. Thanks to its 20 kt LS, JUNO has excellent capabilities to detect <sup>162</sup> all flavour of the  $\mathcal{O}(10 \text{ MeV})$  postshock neutrinos, and using neutrinos of the  $\mathcal{O}(1 \text{ MeV})$  will give <sup>163</sup> informations about the pre-supernovae neutrinos. All those informations will allow to disentangle <sup>164</sup> between the multiple hydro-dynamic models that are currently used to describe the different stage <sup>165</sup> of core-collapse supernovae.

#### <sup>166</sup> **Diffuse supernovae neutrinos background**

<sup>167</sup> Core-collapse supernovae in our galaxy are rare events, but they frequently occur throughout the <sup>168</sup> visible Universe sending burst of neutrinos in direction of the Earth. All those events contributes to <sup>169</sup> a low background flux of low-energy neutrinos called the Diffuse Supernovae Neutrino Background <sup>170</sup> (DSNB). Its flux and spectrum contains informations about the red-shift dependent supernovae rate, <sup>171</sup> the average supernovae neutrino energy and the fraction of black-hole formation in core-collapse <sup>172</sup> supernovae. Depending of the DSNB model, we can expect 2-4 IBD events per year in the energy range <sup>173</sup> above the reactor  $\bar{\nu}_e$  signal, which is competitive with the current Super-Kamiokande+Gadolinium <sup>174</sup> phase [17].

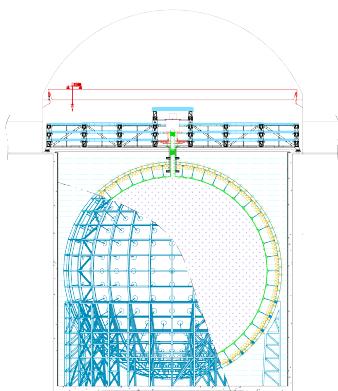
175 **Beyond standard model neutrinos interactions**

176 JUNO will also be able to probe for beyond standard model neutrinos interactions. After the  
177 main physics topics have been accomplished, JUNO could be upgraded to probe for neutrinoless  
178 beta decay ( $0\nu\beta\beta$ ). The detection of such event would give critical informations about the nature  
179 of neutrinos, is it a majorana or a dirac particle. JUNO will also be able to probe for neutrinos that  
180 would come for the decay or annihilation of Dark Matter inside the sun and neutrinos from putative  
181 primordial black hole. Through the unitary test of the mixing matrix, JUNO will be able to search  
182 for light sterile neutrinos. Thanks to JUNO sensitivity, multiple other exotic can be performed on  
183 neutrino related beyond standard model interactions.

184 **2.2 The JUNO detector**

185 The JUNO detector is a scintillator detector buried 693.35 meters under the ground (1800 meters  
186 water equivalent). It consist of Central Detector (CD), a water pool and a Top Tracker (TT) as showed  
187 in figure 2.4a The CD is an acrylic vessel containing the 20 ktons of LS. It is supported by a stainless  
188 steel structure and is immersed in that water pool that is used as shielding from external radiation  
189 and as a cherenkov detector for the background. The top of the experiment is partially covered by  
190 the TT, a plastic scintillator detector which is use to detect the atmospheric muons background acting  
191 as veto detector.

192 The top of the experiment also host the LS purification system, a water purification system, a  
193 ventilation system to get rid of the potential radon in the air. The CD is observed by two system of  
194 Photo-Multipliers Tubes (PMT). They are attached to the steel structure and their electronic readout  
195 is submersed near them. A third system of PMT is also installed on the structure but are facing  
196 outward of the CD, instrumenting the water to be cherenkov detector. The CD and the cherenkov  
197 detector are optically separated by Tyvek sheet. A chimney for LS filling and purification and for  
198 calibration operations connects the CD to the experimental hall from the top.



(A) Schematics view of the JUNO detector.



(B) Top down view of the JUNO detector under construction

199 This section cover in details the different components of the detector and the detection systems.

200 **2.2.1 Principle of detection**

The CD will detect the neutrino and measure their energy mainly via an Inverse Beta Decay (IBD) interaction with proton (mainly  $^{12}\text{C}$  and H) in the LS:

$$\bar{\nu}_e + p \rightarrow n + e^+$$

Kinematics calculation shows that this interaction has an energy threshold for the  $\bar{\nu}_e$  of  $(m_n + m_e - m_p) \approx 1.806$  MeV [18] where  $m_\lambda$  is the mass of the  $\lambda$  particle. This threshold make the experiment blind to very low energy neutrinos. The residual energy  $E_\nu - 1.806$  MeV is be distributed as kinetic energy between the positron and the neutron. The energy of the emitted positron  $E_e$  is given by [18]

$$E_e = \frac{(E_\nu - \delta)(1 + \epsilon_\nu) + \epsilon_\nu \cos \theta \sqrt{(E_\nu - \delta)^2 + \kappa m_e^2}}{\kappa} \quad (2.2)$$

where  $\kappa = (1 + \epsilon_\nu)^2 - \epsilon_\nu^2 \cos^2 \theta \approx 1$ ,  $\epsilon_\nu = \frac{E_\nu}{m_p} \ll 1$  and  $\delta = \frac{m_n^2 - m_p^2 - m_e^2}{2m_p} \ll 1$ . We can see from this equation that the positron energy is strongly correlated to the neutrino energy.

The positron and the neutron will then propagate in the detection medium, the Liquid Scintillator (LS), loosing their kinetic energy by exciting the molecule of the LS (more details in section 2.2.2). Once stopped, the positron will annihilate with an electron from the medium producing two 511 KeV gamma. Those gamma will themselves interact with the LS, exciting it before being absorbed by photoelectrical effect. The neutron will be captured by an hydrogen, emitting a 2.2 MeV gamma in the process. This gamma will also deposit its energy before being absorbed by the LS.

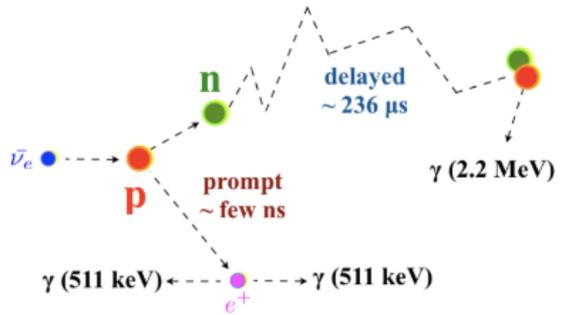


FIGURE 2.5 – Schematics of an IBD interaction in the central detector of JUNO

The scintillation photons have an UV frequency and will be captured by PMTs observing the CD. The analog signal of the PMTs digitized by the electronic is the signal of our experiment. The signal produced by the positron is subsequently called the prompt signal, and the signal coming from the neutron the delayed signal. This naming convention come from the fact that the positron will deposit its energy rather quickly (few ns) where the neutron will take a bit more time ( $\sim 236$   $\mu$ s).

## 2.2.2 Central Detector (CD)

The central detector, composed of 20 ktons of Liquid Scintillator (LS), is the main part of JUNO. The LS is contained in a spherical acrylic vessel supported by a stainless steel structure. The CD and its structural support are submerged in a cylindrical water pool of 43.5m diameter and 44m height. We're confident that the water pool provide sufficient buffer protection in every direction against the rock radioactivity.

### Acrylic vessel

The acrylic vessel is a spherical vessel of inner diameter of 35.4 m and a thickness of 120 mm. It is assembled from 265 acrylic panels, thermo bonded together. The acrylic recipes has been carefully tuned with extensive R&D to ensure it does not include plasticizer and anti-UV material that would stop the scintillation photons. Those panels requires to be pure of radioactive materials to not cause background. Current setup where the acrylic panels are molded in cleanrooms of class 10000, let us reach a uranium and thorium contamination of <0.5 ppt. The molding and thermoforming processes is optimized to increase the assemblage transparency in water to >96%. The acrylic vessel

is supported by a stainless steel structure via supporting node (fig 2.6). The structure and the nodes are designed to be resilient to natural catastrophic events such as earthquake and can support many times the effective load of the acrylic vessel.

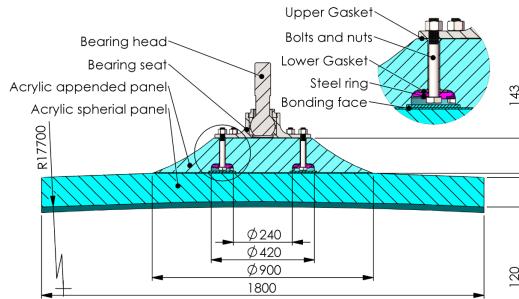


FIGURE 2.6 – Schematics of the supporting node for the acrylic vessel

### 235 Liquid scintillator

236 The Liquid Scintillator (LS) has a similar recipe as the one used in Daya Bay [19] but without  
 237 gadolinium doping. It is made of three components, necessary to shift the wavelength of emitted  
 238 photons to prevent their reabsorption:

- 239 1. The detection medium, the *linear alkylbenzene* (LAB). Selected because of its excellent trans-  
 240 parency, high flash point, low chemical reactivity and good light yield. Accounting for  $\sim$   
 241 98% of the LS, it is the main component with which ionizing particles and gamma interact.  
 242 Charged particles will collide with its electronic cloud transferring energy to the molecules,  
 243 gamma will interact via compton effect with the electronic cloud before finally be absorbed  
 244 via photoelectric effect.
- 245 2. The second component of the LS is the *2,5-diphenyloxazole* (PPO). A fraction of the excitation  
 246 energy of the LAB is transferred to the PPO, mainly via non radiative process [20]. The  
 247 PPO molecules de-excites in the same way, transferring their energy to the bis-MSB. The PPO  
 248 makes for 1.5 % of the LS.
- 249 3. The last component is the *p-bis(o-methylstyryl)-benzene* (bis-MSB). Once excited by the PPO, it  
 250 will emit photon with an average wavelength of  $\sim$  430 nm (full spectrum in figure 2.7) that  
 251 can be detected by our photo-multipliers systems. It amount for  $\sim$  0.5% of the LS.

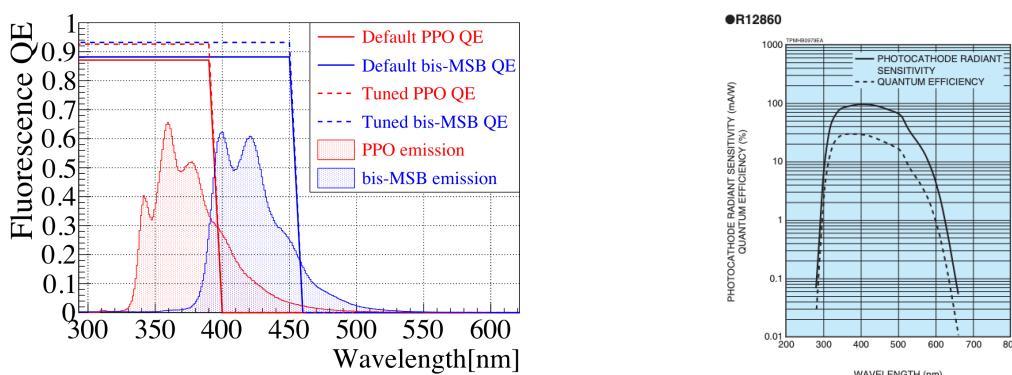


FIGURE 2.7 – On the left: Quantum efficiency (QE) and emission spectrum of the LAB and the bis-MSB [19]. On the right: Sensitivity of the Hamamatsu LPMT depending on the wavelength of the incident photons [21].

This formula has been optimized using dedicated studies with a Daya Bay detector [19, 22] to reach the requirements for the JUNO experiment:

- A light yield / MeV of the amount of  $10^4$  photons to maximize the statistic in the energy measurement.
- An attenuation length comparable to the size of the detector to prevent losing photons during their propagation in the LS. The final attenuation length is 25.8m [23] to compare with the CD diameter of 35.4m.
- Uranium/Thorium radiopurity to prevent background signal. The reactor neutrino program require a contamination fraction  $F < 10^{-15}$  while the solar neutrino program require  $F < 10^{-17}$ .

The LS will frequently be purified and tested in the Online Scintillator Internal Radioactivity Investigation System (OSIRIS) [24] to ensure that the requirements are kept during the lifetime of the experiment.

### Large Photo-Multipliers Tubes(LPMTs)

The scintillation light produced by the LS is then collected by Photo-Multipliers Tubes (PMT) that transform the incoming photon into an electric signal. As described in figure 2.8, the incident photons interact with the photocathode via photoelectric effect producing an electron called a Photo-Electron (PE). This PE is then focused on the dynodes where the high voltage will allow it to be multiplied. After multiple amplification the resulting charge - in coulomb [C] - is collected by the anode and the resulting electric signal can be digitalized by the readout electronics from which the charge and timing can be extracted.

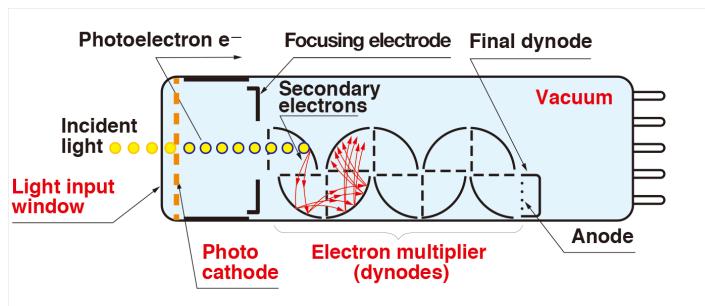


FIGURE 2.8 – Schematic of a PMT

The Large Photo-Multipliers Tubes (LPMT), used in the central detector and in the water pool, are 20-inch (50.8 cm) radius PMTs.  $\sim 5000$  dynode-PMTs [21] were produced by the Hamamatsu<sup>®</sup> company and  $\sim 15000$  Micro-Channel Plate (MCP) [25] by the NNVT<sup>®</sup> company. This system is the one responsible for the energy measurement with a energy resolution of  $3\%/\sqrt{E}$ , resolution necessary for the mass ordering measurement. To reach this precision, the system is composed of 17612 PMTs quasi uniformly distributed over the detector for a coverage of 75.2% reaching  $\sim 1800$  PE/MeV or  $\sim 2.3\%$  resolution due to statistic, leaving  $\sim 0.7\%$  for the systematic uncertainties. To maintain the resolution over the lifetime of the experiment, JUNO require a failure rate  $< 1\%$  over 6 years.

The LPMTs electronic are divided in two parts. One "near", located underwater, in proximity of the LPMT to reduce the cable length between the PMT and early electronic. A second one, outside of the detector that is responsible for higher level analysis before sending the data to the DAQ.

The light yield per MeV induce that a LPMT can collect between 1 and 1000 PE per event, causing non linearity in the PMT response that need to be understood and calibrated, see section 2.3 for more details.

288 **Small Photo-Multipliers Tubes (SPMTs)**

289 The Small PMT (SPMTs) system is made of 3-inch (7.62 cm) PMTs. They will be used in the CD  
 290 as a secondary detection system. Those 25600 SPMTs will observe the same events as the LPMTs,  
 291 thus sharing the physics and detector systematics up until the photon conversion. With a detector  
 292 coverage of 2.7%, this system will collect  $\sim 43$  PE/MeV for a final energy resolution of  $\sim 17\%$ .  
 293 This resolution is not enough to measure the NMO,  $\theta_{13}$ ,  $\Delta m_{31}^2$  but will be sufficient to independently  
 294 measure  $\theta_{12}$  and  $\Delta m_{21}^2$ .

295 Due to the low PE rate, SPMTs will be running in photo-counting mode in the reactor range thus  
 296 will be insensitive to non-linearity effect. Also, due to their smaller size and electronics, SPMTs have  
 297 a better timing resolutions than the LPMTs. At higher energy range, like supernovae events, LPMTs  
 298 will saturate where SPMTs due to their lower PE collection will to produce a reliable measure of the  
 299 energy spectrum.

300 The Data Acquisition System (DAQ) is designed to support the event rate of IBD, background,  
 301 dark noise and supplementary storage buffers are present in the LPMT electronics to withstand the  
 302 event rate during supernovae burst.

303 **2.2.3 Veto detector**

304 The CD will be bathed in constant background noise coming from numerous sources : the ra-  
 305 dioactivity from surrounding rock and its own components or from the flux of cosmic muons. This  
 306 background needs to be rejected to ensure the purity of the IBD spectrum. To prevent a big part of  
 307 them, JUNO use two veto detector that will tag events as background before CD analysis.

308 **Cherenkov in water pool**

309 The Water Cherenkov Detector (WCD) is the instrumentalization of the water buffer around the  
 310 CD. When high speed charged particles will pass through the water, they will produced cherenkov  
 311 photons. The light will be collected by 2400 MCP LPMTs installed on the outer surface of the CD  
 312 structure. The muons veto strategy is based on a PMT multiplicity condition. WCD PMTs are  
 313 grouped in ten zones: 5 in the top, 5 in the bottom. A veto is raised either when more than 19  
 314 PMTs are triggered in one zone or when two adjacent zones simultaneously trigger more than 13  
 315 PMTs. Using this trigger, we expect to reach a muon detection efficiency of 99.5% while keeping the  
 316 noise at reasonable level.

317 **Top tracker**

318 The JUNO Top Tracker (TT) is a plastic scintillator detector located on the top of the experiment  
 319 (see figure 2.9). Made from plastic scintillator from OPERA [26] layered horizontally in 3 layers on  
 320 the top of the detector, the TT will be able to detect incoming atmospheric muons. With its coverage,  
 321 about 1/3 of the of all atmospheric muons that passing through the CD will also pass through the 3  
 322 layer of the detector. While it does not cover the majority of the CD, the TT is particularly effective  
 323 to detect muons coming through the filling chimney region which might present difficulties from the  
 324 other subsystems in some classes of events.

325 **2.3 Calibration strategy**

326 The calibration is a crucial part of the JUNO experiment. Because we are looking at civil reactor  
 327 neutrino it might be impossible to run measurement without signal, it would need to shut down  
 328 every reactor from the Taishan and Yangjiang power plants which is realistically impossible. Because  
 329 of this continuous rate, low frequency signal event, we need high frequency, recognisable sources in  
 330 the energy range of interest, [0-12] MeV for the positron signal and 2.2 MeV for the neutron capture.  
 331 It is understood that the CD response will be different depending on the type of particle, due to the

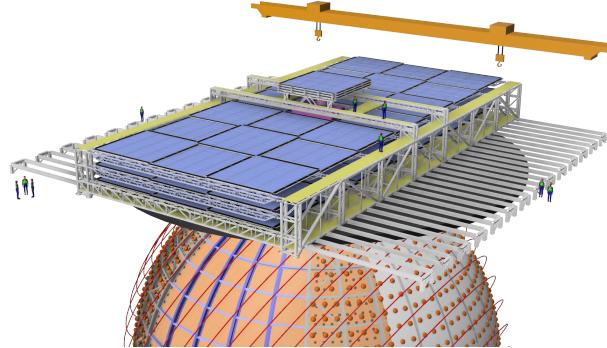


FIGURE 2.9 – The JUNO top tracker

interaction with LS, and on the position on the event, due to the absorption length of the LS and the optical response of the acrylic sphere (see section 2.5). The energy response is also expected to be non-linear due to the LS response [19] but also due to the saturation of the LPMTs system when collecting a large amount of PE [27].

### 2.3.1 Energy scale calibration

While electrons and positrons sources would be ideal, for a large LS detector thin-walled electrons or positrons sources could lead to leakage of radionuclides causing contamination. Instead are considered gamma sources in the range of the prompt energy of IBDs. The sources are reported in table 2.4.

| Sources / Processes             | Type        | Radiation                                |
|---------------------------------|-------------|--|
| $^{137}\text{Cs}$               | $\gamma$    | 0.0662 MeV                               |
| $^{54}\text{Mn}$                | $\gamma$    | 0.835 MeV                                |
| $^{60}\text{Co}$                | $\gamma$    | 1.173 + 1.333 MeV                        |
| $^{40}\text{K}$                 | $\gamma$    | 1.461 MeV                                |
| $^{68}\text{Ge}$                | $e^+$       | annihilation 0.511 + 0.511 MeV           |
| $^{241}\text{Am-Be}$            | $n, \gamma$ | neutron + 4.43 MeV ( $^{12}\text{C}^*$ ) |
| $^{241}\text{Am-}^{13}\text{C}$ | $n, \gamma$ | neutron + 6.13 MeV ( $^{16}\text{O}^*$ ) |
| $(n, \gamma)p$                  | $\gamma$    | 2.22 MeV                                 |
| $(n, \gamma)^{12}\text{C}$      | $\gamma$    | 4.94 MeV or 3.68 + 1.26 MeV              |

TABLE 2.4 – List of sources and their process considered for the energy scale calibration

For the  $^{68}\text{Ge}$  source, it'll decay in  $^{68}\text{Ga}$  via electron capture, which will itself  $\beta^+$  decay into  $^{68}\text{Zn}$ . The positrons will be absorbed by the enclosure so only the annihilation gamma will be released. In addition,  $(\alpha, n)$  sources like  $^{241}\text{Am-Be}$  and  $^{241}\text{Am-}^{13}\text{C}$  are used to provide both high energy gamma and neutrons, which will later be captured in the LS producing the 2.2 MeV gamma.

From this calibration we call  $E_{vis}$  the "visible energy" that is reconstructed by our current algorithms and we compare it to the true energy deposited by the calibration source. The results shown in figure 2.10 show the expected response of the detector from calibration sources. The non-linearity is clearly visible from the  $E_{vis}/E_{true}$  shape. See [28] for more details.

### 2.3.2 Calibration system

The JUNO detector will need to investigate the non-uniformity in the detector response depending of the position of the event (more details in section 2.5). For this we will able to deploy sources

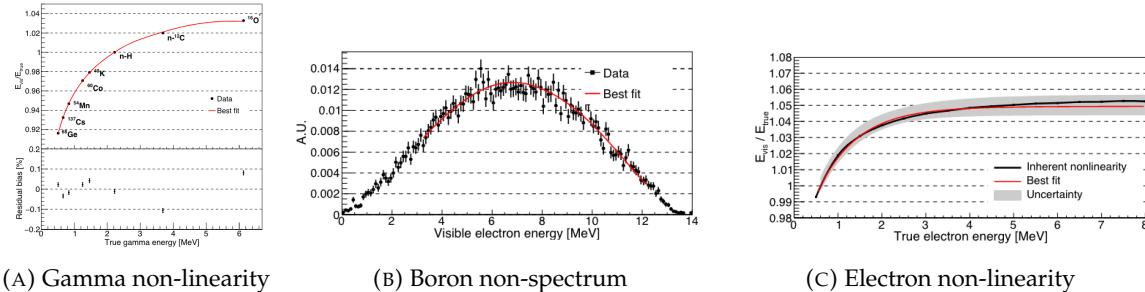


FIGURE 2.10 – Fitted and simulated non linearity of gamma and electron sources and  $^{12}\text{B}$  spectrum. Black points are simulated data while red points are the best fit

352 using multiples systems that are schematized in figure 2.11 :

- 353 — For a one-dimension vertical calibration, the Automatic Calibration Unit (ACU) will be able  
354 to deploy multiple radioactive sources or a pulse laser diffuser ball along the central axis of  
355 the CD through the top chimney. The source position precision is less than 1cm.
- 356 — For off-axis calibration, a calibration source attached to a Cable Loop System (CLS) can be  
357 moved on a vertical half-plane by adjusting the length of two connection cable. Two set of  
358 CSL will be deployed to provide a 79% effective coverage of a vertical plane.
- 359 — A Guiding Tube (GT) will surround the CD to calibrate the non-uniformity of the response at  
360 the edge of the detector
- 361 — A Remotely Operated under-LS Vehicle (ROV) can be deployed to desired location inside LS  
362 for a more precise and comprehensive calibration. The ROV will also be equipped with a  
363 camera for inspection of the CD.

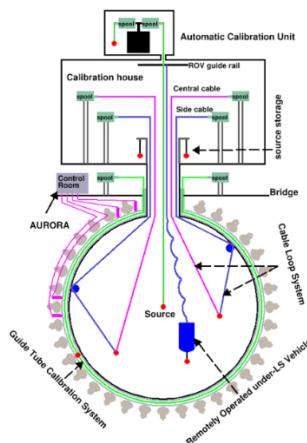


FIGURE 2.11 – Overview of the calibration system

- 364 The preliminary calibration program is depicted in table 2.5.

## 2.4 Software

366 The simulation, reconstruction and analysis algorithms are all packaged in the JUNO software,  
367 subsequently called the software. It is composed of multiple components integrated in the SNiPER  
368 [29] framework:

- 369 — Various primary particles simulators for the different kind of events, background and calibra-  
370 tion sources.

| Program                   | Purpose           | System          | Duration [min] |
|---------------------------|-------------------|-----------------|----------------|
| Weekly calibration        | Neutron (Am-C)    | ACU             | 63             |
|                           | Laser             | ACU             | 78             |
| Monthly calibration       | Neutron (Am-C)    | ACU             | 120            |
|                           | Laser             | ACU             | 147            |
|                           | Neutron (Am-C)    | CLS             | 333            |
|                           | Neutron (Am-C)    | GT              | 73             |
| Comprehensive calibration | Neutron (Am-C)    | ACU, CLS and GT | 1942           |
|                           | Neutron (Am-Be)   | ACU             | 75             |
|                           | Laser             | ACU             | 391            |
|                           | $^{68}\text{Ge}$  | ACU             | 75             |
|                           | $^{137}\text{Cs}$ | ACU             | 75             |
|                           | $^{54}\text{Mn}$  | ACU             | 75             |
|                           | $^{60}\text{Co}$  | ACU             | 75             |
|                           | $^{40}\text{K}$   | ACU             | 158            |

TABLE 2.5 – Calibration program of the JUNO experiment

- A Geant4 [30–32] Monte Carlo (MC) simulation simulating the detectors geometries, a custom optical model for the LS and the supporting structures of the detectors. The Geant4 simulation integrate all relevant physics process for JUNO, validated by the collaboration. This step of the simulation is commonly called *Detsim* and compute up to the production of photo-electrons in the PMTs. The optics properties of the different materials and detector components have been measured beforehand to be used to define the material and surfaces in the simulation
  - An electronic simulation, simulating the response waveform of the PMTs, tracking it through the digitization process, accounting for effects such as non-linearity, dark noise, Time Transit Spread (TTS), pre-pulsing, after-pulsing and ringing if the waveform. This step is commonly referenced as *Elecsim*.
  - A waveform reconstruction where the digitized waveform are filtered to remove high-frequency white noise and then deconvoluted to yield time and charge informations of the photons hits on the PMTs. This step is commonly referenced as *Calib*.
  - The charge and time informations are used by reconstruction algorithms to reconstruct the interaction vertex and the deposited energy. This step is commonly reported as *Reco*. See section 2.5 for more details on the reconstruction.
  - Once the singular events are reconstructed, they go through event pairing and classification to select IBD events. This step is named Event Classification.
  - The purified signal is then analysed by the analysis framework which depend of the physics topic of interest.
- The steps Reco and Event Classification are divided into two category of algorithm. Fast but less accurate algorithms that are running during the data taking designed as *Online* algorithms. Those algorithm are used to take the decision to save the event on tape or to throw it away. More accurate algorithms that run on batch of events designed *Offline* algorithms. They are used for the physics analysis and the Offline Reco will be a topic of interest for this thesis.

## 2.5 State of the art of the Offline IBD reconstruction in JUNO

The main reconstruction method currently run in JUNO is a data-driven method based on a likelihood maximization [33, 34] using only the LPMTs. The first step is to reconstruct the interaction vertex from which the energy reconstruction is dependent. It is also necessary for event pairing and classification.

401 **2.5.1 Interaction vertex reconstruction**

402 To start the likelihood maximization, a rough estimation of the vertex and of the event timing is  
 403 needed. We start by estimating the vertex position using a charge based algorithm.

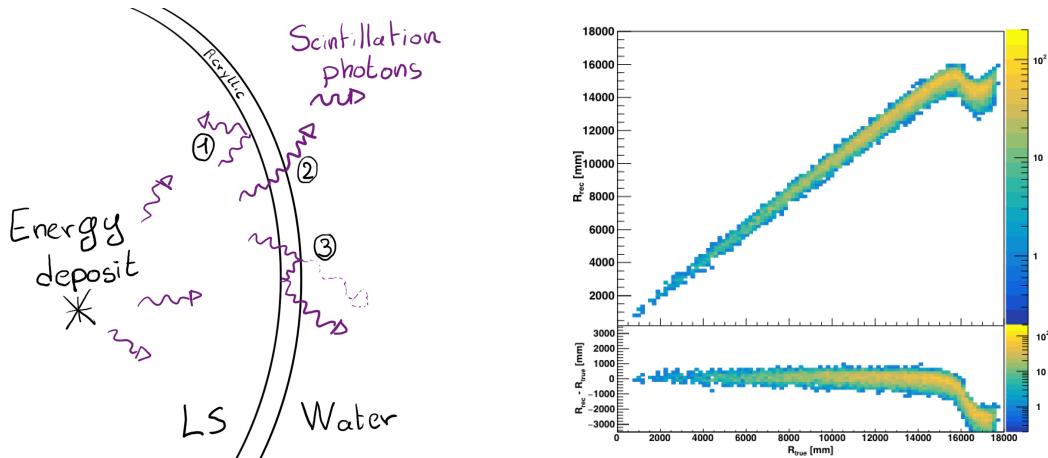
404 **Charge based algorithm**

405 The charge-based algorithm is basically base on the charge-weighted average of the PMT position.

$$\vec{r}_{cb} = a \cdot \frac{\sum_i q_i \cdot \vec{r}_i}{\sum_i q_i} \quad (2.3)$$

406 Where  $q_i$  is the reconstructed charge of the pulse of the  $i$ th PMT and  $\vec{r}_i$  is its position.  $\vec{r}_0$  is the  
 407 reconstructed interaction position.  $a$  is a scale factor introduced because a weighted average over  
 408 a 3D sphere is inherently biased. Using calibration we can estimate  $a \approx 1.3$  [35]. The results in  
 409 figure 2.12b shows that the reconstruction is biased from around 15m and further. This is due to the  
 410 phenomena called “total reflection area” or TR Area.

411 As depicted in the figure 2.12a the optical photons, given that they have a sufficiently large  
 412 incidence angle, can be deviated of their trajectories when passing through the interfaces LS-acrylic  
 413 and water-acrylic due to the optical index difference. This cause photons to be lost or to be detected  
 414 by PMT further than anticipated if we consider their rectilinear trajectories. This cause the charge  
 415 barycenter the be located closer to the center than the vent really is.



(A) Illustration of the different optical photons reflection scenarios. 1 is the reflection of the photon at the interface LS-acrylic or acrylic-water. 2 is the transmission of the photons through the interfaces. 3 is the conduction of the photon in the acrylic.

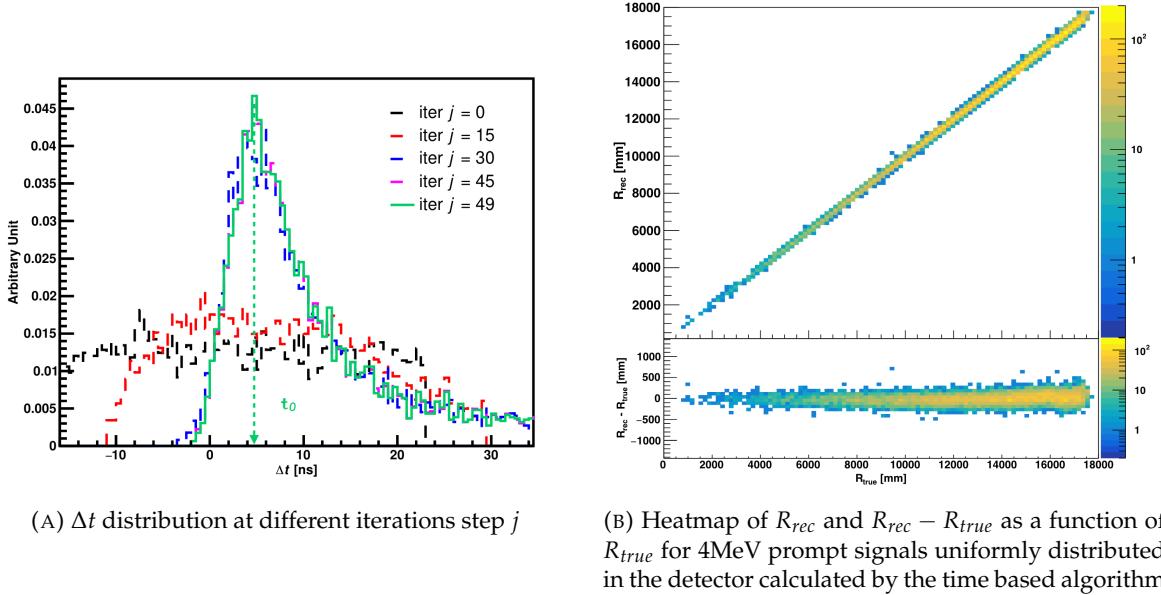
(B) Heatmap of  $R_{rec}$  and  $R_{rec} - R_{true}$  as a function of  $R_{true}$  for 4MeV prompt signals uniformly distributed in the detector calculated by the charge based algorithm

FIGURE 2.12

416 It is to be noted that charge based algorithm, in addition to be biased near the edge of the detector,  
 417 does not provide any information about the timing of the event. Therefore, a time based algorithm  
 418 needs to be introduced to provide initial values.

419 **Time based algorithm**

420 The time based algorithm use the distribution of the time of flight corrections  $\Delta t$  (Eq 2.4) of an  
 421 event to reconstruct its vertex and  $t_0$ . It follow the following iterations:

(A)  $\Delta t$  distribution at different iterations step  $j$ (B) Heatmap of  $R_{rec}$  and  $R_{rec} - R_{true}$  as a function of  $R_{true}$  for 4MeV prompt signals uniformly distributed in the detector calculated by the time based algorithm

- 422 1. Use the charge based algorithm to get an initial vertex to start the iteration.  
 423 2. Calculate the time of flight correction for the  $i$ th PMT using

$$\Delta t_i(j) = t_i - \text{tof}_i(j) \quad (2.4)$$

424 where  $j$  is the iteration step,  $t_i$  is the timing of the  $i$ th PMT, and  $\text{tof}_i$  is the time-of-flight of the  
 425 photon considering an rectilinear trajectory and an effective velocity in the LS and water (see  
 426 [35] for detailed description of this effective velocity). Plot the  $\Delta t$  distribution and label the  
 427 peak position as  $\Delta t^{\text{peak}}$  (see fig 2.13a).

- 428 3. Calculate a correction vector  $\vec{\delta}[\vec{r}(j)]$  as

$$\vec{\delta}[\vec{r}(j)] = \frac{\sum_i \left( \frac{\Delta t(j) - \Delta t^{\text{peak}}(j)}{\text{tof}_i(j)} \right) \cdot (\vec{r}_0(j) - \vec{r}_i)}{N^{\text{peak}}(j)} \quad (2.5)$$

429 where  $\vec{r}_0$  is the vertex position at the beginning of this iteration,  $\vec{r}_i$  is the position of the  $i$ th  
 430 PMT. To minimize the effect of scattering, dark noise and reflection, only the pulse happening  
 431 in a time window (-10 ns, +5 ns) around  $\Delta t^{\text{peak}}$  are considered.  $N^{\text{peak}}$  is the number of PE  
 432 collected in this time-window.

- 433 4. if  $\vec{\delta}[\vec{r}(j)] < 1\text{mm}$  or  $j \geq 100$ , stop the iteration. Otherwise  $\vec{r}_0(j+1) = \vec{r}_0(j) + \vec{\delta}[\vec{r}(j)]$  and go to  
 434 step 2.

435 However because the earliest arrival time is used,  $t_i$  is related to the number photoelectrons  $N_i^{\text{pe}}$   
 436 detected by the PMT [36–38]. To reduce bias in the vertex reconstruction, the following equation is  
 437 used to correct  $t_i$  into  $t'_i$ :

$$t'_i = t_i - p_0 / \sqrt{N_i^{\text{pe}}} - p_1 - p_2 / N_i^{\text{pe}} \quad (2.6)$$

438 The parameters  $(p_0, p_1, p_2)$  were optimized to (9.42, 0.74, -4.60) for Hamamatsu PMTs and (41.31,  
 439 -12.04, -20.02) for NNVT PMTs [35]. The results presented in figure 2.13b shows that the time based  
 440 algorithm provide a more accurate vertex and is unbiased even in the TR area. This results  $(\vec{r}_0, t_0)$  is  
 441 used as initial value for the likelihood algorithm.

442 **Time likelihood algorithm**

443 The time likelihood algorithm use the residual time expressed as follow

$$t_{\text{res}}^i(\vec{r}_0, t_0) = t_i - \text{tof}_i - t_0 \quad (2.7)$$

444 In a first order approximation, the scintillator time response Probability Density Function (PDF)  
 445 can be described as the emission time profile of the scintillation photons, the Time Transit Spread  
 446 (TTS) and the dark noise of the PMTs. The emission time profile  $f(t_{\text{res}})$  is described like

$$f(t_{\text{res}}) = \sum_k \frac{\rho_k}{\tau_k} e^{-\frac{t_{\text{res}}}{\tau_k}}, \sum_k \rho_k = 1 \quad (2.8)$$

447 as the sum of the  $k$  component that emit light in the LS each one characterised by it's decay time  $\tau_k$   
 448 and intensity fraction  $\rho_k$ . The TTS component is expressed as a gaussian convolution

$$g(t_{\text{res}}) = \frac{1}{\sqrt{2\pi}\sigma} e^{-\frac{(t_{\text{res}}-\nu)^2}{2\sigma^2}} \cdot f(t_{\text{res}}) \quad (2.9)$$

449 where  $\sigma$  is the TTS of PMTs and  $\nu$  is the average transit time. The dark noise is not correlated with any  
 450 physical events and considered as constant rate over the time window considered  $T$ . By normalizing  
 451 the dark noise probability  $\epsilon(t_{\text{res}})$  as  $\int_T \epsilon(t_{\text{res}}) dt_{\text{res}} = \epsilon_{\text{dn}}$ , it can be integrated in the PDF as

$$p(t_{\text{res}}) = (1 - \epsilon_{\text{dn}}) \cdot g(t_{\text{res}}) + \epsilon(t_{\text{res}}) \quad (2.10)$$

452 The distribution of the residual time  $t_{\text{res}}$  of an event can then be compared to  $p(t_{\text{res}})$  and the best  
 453 vertex  $\vec{r}_0$  and  $t_0$  can be chosen by minimizing

$$\mathcal{L}(\vec{r}_0, t_0) = -\ln \left( \prod_i p(t_{\text{res}}^i) \right) \quad (2.11)$$

454 The parameter of Eq. 2.10 can be measured experimentally. The results shown in figure 2.14  
 455 used PDF from monte carlo simulation. The results shows that  $R_{\text{rec}} - R_{\text{true}}$  is biased depending  
 456 on the energy. While this could be corrected using calibration, another algorithm based on charge  
 457 likelihood was developed to correct this problem.

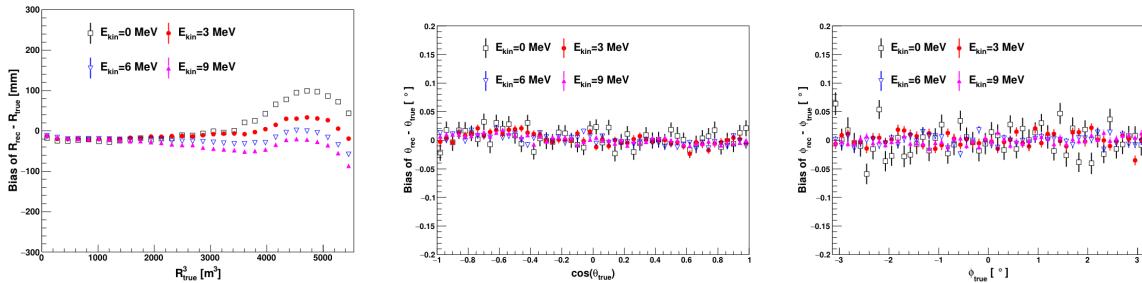


FIGURE 2.14 – Bias of the reconstructed radius  $R$  (left),  $\theta$  (middle) and  $\phi$  (right) for multiple energies by the time likelihood algorithm

458 **Charge likelihood algorithm**

459 Similarly to the time likelihood algorithms that use a timing PDF, the charge likelihood algorithm  
 460 use PE PDF in each PMT depending on the energy and position of the event. With  $\mu(\vec{r}_0, E)$  the mean

expected number of PE detected by each PMT, the probability to observe  $N_{pe}$  in a PMT follow a Poisson distribution. Thus

— The probability to observe no hit ( $N_{pe} = 0$ ) in the  $j$ th PMT is  $P_{nohit}^j(\vec{r}_0, E) = e^{-\mu_j}$

— The probability to observe  $N_{pe} \neq 0$  in the  $i$ th PMT is  $P_{hit}^i(\vec{r}_0, E) = \frac{\mu^{N_{pe}} e^{-\mu_i}}{N_{pe}^i!}$

Therefore, the probability to observe a specific hit pattern can be expressed as

$$P(\vec{r}_0, E) = \prod_j P_{nohit}^j(\vec{r}_0, E) \cdot \prod_i P_{hit}^i(\vec{r}_0, E) \quad (2.12)$$

The best fit values of  $\vec{R}_0$  and  $E$  can then be calculated by minimizing the negative log-likelihood

$$\mathcal{L}(\vec{r}_0, E) = -\ln(P(\vec{r}_0, E)) \quad (2.13)$$

In principle,  $\mu_i(\vec{r}_0, E)$  could be expressed

$$\mu_i(\vec{r}_0, E) = Y \cdot \frac{\Omega(\vec{r}_0, r_i)}{4\pi} \cdot \epsilon_i \cdot f(\theta_i) \cdot e^{-\sum_m \frac{d_m}{\zeta_m}} \cdot E + \delta_i \quad (2.14)$$

where  $Y$  is the energy scale factor,  $\Omega(\vec{r}_0, r_i)$  is the solid angle of the  $i$ th PMT,  $\epsilon_i$  is its detection efficiency,  $f(\theta_i)$  its angular response,  $\zeta_m$  is the attenuation length in the materials and  $\delta_i$  the expected number of dark noise.

However Eq. 2.14 assume that the scintillation light yield is linear with energy and describe poorly the contribution of indirect light, shadow effect due to the supporting structure and the total reflection effects. The solution is to use data driven methods to produce the pdf by using the calibrations sources and position described in section 2.3. In the results presented in figures 2.15, the PDF was produced using MC simulation and 29 specific calibrations position [35] along the Z-axis of the detector. We see that the charge likelihood algorithm show little bias in the TR area and a

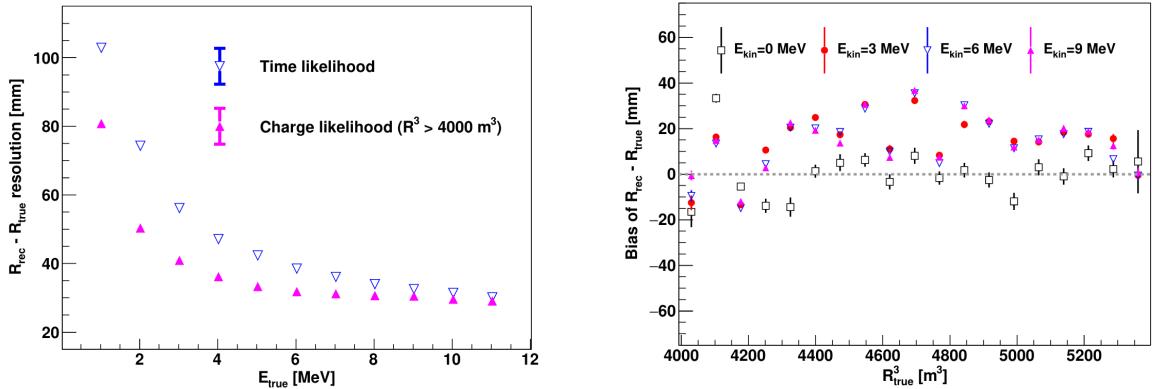


FIGURE 2.15 – On the left: Resolution of the reconstructed  $R$  as a function of the energy in the TR area ( $R^3 > 4000 \text{ m}^3 \equiv R > 16 \text{ m}$ ) by the charge and time likelihood algorithms. On the right: Bias of the reconstructed  $R$  in the TR area for different energies by the charge likelihood algorithm

better resolution than the time likelihood. The figure 2.16 shows the radial resolution of the different algorithm presented for this section, we can see the refinement at each step and that the charge likelihood yield the best results.

The charge based likelihood algorithms already give some information on the energy as Eq. 2.13 is minimized but the energy can be further refined as shown in the next section.

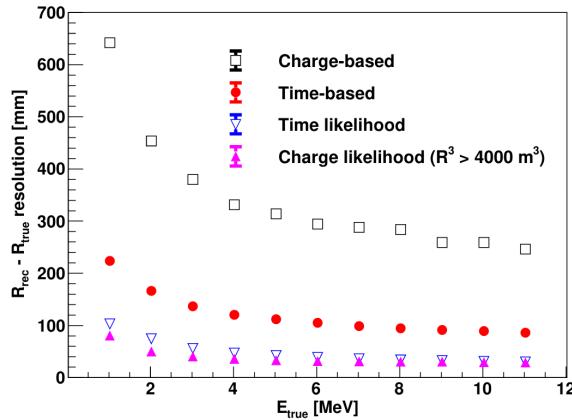


FIGURE 2.16 – Radial resolution of the different vertex reconstruction algorithms as a function of the energy

### 482 2.5.2 Energy reconstruction

483 As explained in section 2.1.1, energy resolution is crucial for the NMO and oscillation parameters  
 484 measurements. Thus the energy reconstruction algorithm should take into consideration as much  
 485 detector effect as possible. The following method is a data driven method based on calibrations  
 486 sample inspired by the charge likelihood algorithm described above [39].

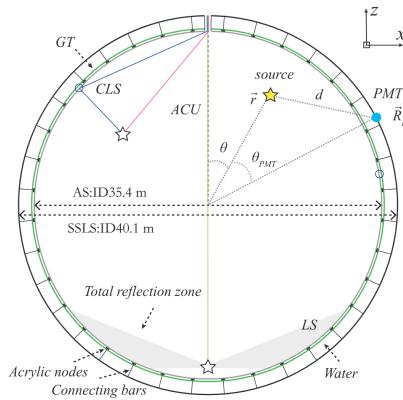


FIGURE 2.17 – Definition of the variables used in the energy reconstruction

### 487 Charge estimation

488 The most important element in the energy reconstruction is  $\mu_i(\vec{r}_0, E)$  described in Eq. 2.14. For  
 489 realistic cases, we also need to take into account the electronics effect that were omitted in the  
 490 previous section. Those effect will cause a charge smearing due to the uncertainties in the  $N_{pe}$   
 491 reconstruction. Thus we define  $\hat{\mu}^L(\vec{r}_0, E)$  which is the expected  $N_{pe}/E$  in the whole detector for an  
 492 event with visible energy  $E_{vis}$  and position  $\vec{r}_0$ . The position of the event and PMTs are now defined

<sup>493</sup> using  $(r, \theta, \theta_{pmt})$  as defined in figure 2.17.

$$\hat{\mu}(r, \theta, \theta_{pmt}, E_{vis}) = \frac{1}{E_{vis}} \frac{1}{M} \sum_i^M \frac{\bar{q}_i - \mu_i^D}{\text{DE}_i}, \quad \mu_i^D = \text{DNR}_i \cdot L \quad (2.15)$$

<sup>494</sup> where  $i$  runs over the PMTs with the same  $\theta_{pmt}$ ,  $\text{DE}_i$  is the detection efficiency of the  $i$ th PMT.  $\mu_i^D$   
<sup>495</sup> is the expected number of dark noise photoelectrons in the time window  $L$ . The time window have  
<sup>496</sup> been optimized to  $L = 280$  ns [39].  $\bar{q}_i$  is the average recorded photoelectrons in the time window  
<sup>497</sup> and  $\hat{Q}_i$  is the expected average charge for 1 photoelectron. The  $N_{pe}$  map is constructed following the  
<sup>498</sup> procedure described in [34].

#### <sup>499</sup> Time estimation

<sup>500</sup> The second important observable is the hit time of photons that was previously defined in Eq.  
<sup>501</sup> 2.7. It is here refined as

$$t_r = t_h - \text{tof} - t_0 = t_{LS} + t_{TT} \quad (2.16)$$

<sup>502</sup> where  $t_h$  is the time of hit,  $t_{LS}$  is the scintillation time and  $t_{TT}$  the transit time of PMTs that is described  
<sup>503</sup> by a gaussian

$$t_{TT} = \mathcal{N}(\bar{\mu}_{TT} + t_d, \sigma_{TT}) \quad (2.17)$$

<sup>504</sup> where  $\mu_{TT}$  is the mean transit time in PMTs,  $\sigma_{TT}$  is the Transit Time Spread (TTS) of the PMTs and  $t_d$   
<sup>505</sup> is the delay time in the electronics. The effective refraction index of the LS is also corrected to take  
<sup>506</sup> into account the propagation distance in the detector.

<sup>507</sup> The timing PDF  $P_T(t_r | r, d, \mu_l, \mu_d, k)$  can now be generated using calibration sources [39]. This PDF  
<sup>508</sup> describe the probability that the residual time of the first photon hit is in  $[t_r, t_r + \delta]$  with  $r$  the radius  
<sup>509</sup> of the event vertex,  $d = |\vec{r} - \vec{r}_{PMT}|$  the propagation distance,  $\mu_l$  and  $\mu_d$  the expected number of PE  
<sup>510</sup> and dark noise in the electronic reading window and  $k$  is the detected number of PE.

<sup>511</sup> Now let denote  $f(t, r, d)$  the probability density function of "photoelectron hit a time t" for an  
<sup>512</sup> event happening at  $r$  where the photons traveled the distance  $d$  in the LS

$$F(t, r, d) = \int_t^L f(t', r, d) dt' \quad (2.18)$$

<sup>513</sup> Based on the PDF for one photon  $k = 1$ , one can define

$$P_T^l(t | k = n) = I_n^l [f_l(t) F_l^{n-1}(t)] \quad (2.19)$$

<sup>514</sup> where the indicator  $l$  means that the photons comes from the LS and  $I_n^l$  a normalisation factor. To this  
<sup>515</sup> pdf we add the probability to have photons coming from the dark noise indicated by the indicator  $d$   
<sup>516</sup> using

$$f_d(t) = 1/L, \quad F_d(t) = 1 - \frac{t}{L} \quad (2.20)$$

<sup>517</sup> and so for the case where only one photon is detected by the PMT ( $k = 1$ )

$$P_T(t | \mu_l, \mu_d, k = 1) = I_1 [P(1, \mu_l) P(0, \mu_d) f_l(t) + P(0, \mu_l) P(1, \mu_d) f_d(t)] \quad (2.21)$$

<sup>518</sup> where  $P(k_\alpha, \mu_\alpha)$  is the Poisson probability to detect  $k_\alpha$  PE from  $\alpha \in \{l, d\}$  with the condition  $k_l + k_d =$   
<sup>519</sup>  $k$ .

<sup>520</sup> Now that we have the individual timing and charge probability we can construct the charge  
<sup>521</sup> likelihood referred as QMLE:

$$\mathcal{L}(q_1, q_2, \dots, q_N | \vec{r}, E_{vis}) = \prod_{j \in \text{unfired}} e^{-\mu_j} \prod_{i \in \text{fired}} \left( \sum_{k=1} P_Q(q_i | k) \cdot P(k, \mu_i) \right) \quad (2.22)$$

522 where  $\mu_i = E_{vis}\hat{\mu}_i^L + \mu_i^D$  and  $P(k, \mu_i)$  is the Poisson probability of observing  $k$  PE.  $P_Q(q_i|k)$  is the  
 523 charge pdf for  $k$  PE. And we can also construct the time likelihood referred as TMLE:

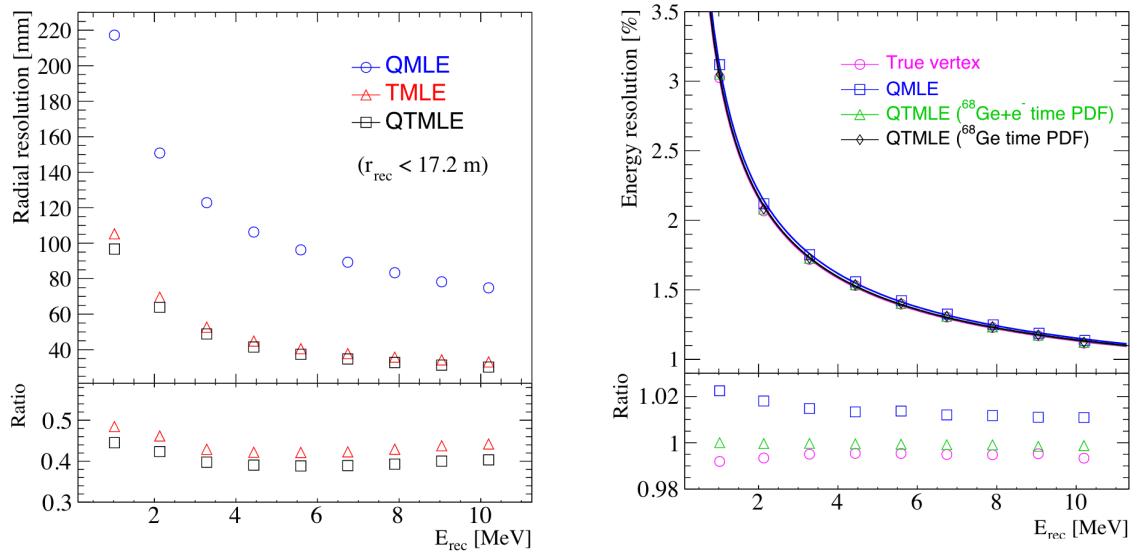
$$\mathcal{L}(t_{1,r}, t_{2,r}, \dots, t_{N,r} | \vec{r}, t_0) = \prod_{i \in \text{hit}} \frac{\sum_{k=1}^K P_T(t_{i,r} | r, d, \mu_i^l, \mu_i^d, k) \cdot P(k, \mu_i^l + \mu_i^d)}{\sum_{k=1}^K P(k, \mu_i^l + \mu_i^d)} \quad (2.23)$$

524 where  $K$  is cut to 20 PE and hit is the set of hits satisfying  $-100 < t_{i,r} < 500$  ns.

525 Merging those two likelihood give the charge-time likelihood QTMLLE

$$\mathcal{L}(q_1, q_2, \dots, q_N; t_{1,r}, t_{2,r}, \dots, t_{N,r} | \vec{r}, t_0, E_{vis}) = \mathcal{L}(q_1, q_2, \dots, q_N | \vec{r}, E_{vis}) \cdot \mathcal{L}(t_{1,r}, t_{2,r}, \dots, t_{N,r} | \vec{r}, t_0) \quad (2.24)$$

526 The radial and energy resolutions of the different likelihood are presented in figure 2.18 (from  
 527 [39]). We can see the improvement of adding the time information to the vertex reconstruction and  
 528 that an increase in vertex precision can bring improvement in the energy resolution, especially at low  
 529 energies.



(A) Radial resolutions of the likelihood-based algorithm TMLE, QMLE and QTMLLE

(B) Energy resolution of QMLE and QTMLLE using different vertex resolutions

FIGURE 2.18

530 Data driven methods prove to be performant in the energy and vertex reconstruction given that  
 531 we have enough calibrations sources to produce the PDF. In the next section, we'll see another type  
 532 of data-driven method based on machine learning.

### 533 2.5.3 Machine learning for reconstruction

534 Machine learning (ML) is family of data-driven algorithms that are inferring behavior and results  
 535 from a training dataset. A overview of methods and detailed explanation of the Neural Network  
 536 (NN) subfamily can be found in Chapter 3.

537   **Vertex reconstruction**

538   **Energy reconstruction**

539   **2.6 JUNO sensitivity to NMO and precise measurements**

540   **2.6.1 Fitting procedure**



<sup>541</sup> **Chapter 3**

<sup>542</sup> **Machine learning and Artificial  
<sup>543</sup> Neural Network**

<sup>544</sup> **3.1 History of the Machine learning**

<sup>545</sup> **3.2 Boosted Decision Tree (BDT)**

<sup>546</sup> **3.3 Artificial Neural Network (NN)**

<sup>547</sup> **3.3.1 Fully Connected Deep Neural Network (FCDNN)**

<sup>548</sup> **3.3.2 Convolutional Neural Network (CNN)**

<sup>549</sup> **3.3.3 Graph Neural Network (GNN)**

<sup>550</sup> **3.3.4 Adversarial Neural Network (ANN)**

<sup>551</sup> **Generative Adversarial Network (GAN)**

<sup>552</sup> **Reinforcement Learning (RL)**

<sup>553</sup> **Random Search (RS)**

<sup>554</sup> **Bayesian Optimization**



<sup>555</sup> **Chapter 4**

<sup>556</sup> **Image recognition for IBD  
reconstruction with the SPMT system**

<sup>557</sup>



<sup>558</sup> **Chapter 5**

<sup>559</sup> **Graph representation of JUNO for IBD  
reconstruction with the LPMT system**

<sup>560</sup>



<sup>561</sup> Chapter 6

<sup>562</sup> **Reliability of machine learning  
methods**

<sup>563</sup>



<sup>564</sup> **Chapter 7**

<sup>565</sup> **Discrimination of e+/e- events in  
JUNO**

<sup>566</sup>



<sup>567</sup> **Chapter 8**

<sup>568</sup> **Conclusion**



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