

Binary black hole detections from LIGO-VIRGO runs 1 and 2

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12/04/2023

Abstract

Introduction and Background

Gravitational waves as first predicted by Albert Einstein in 1915 in his paper on special and general relativity, are ripples in the fabric of space-time due to the acceleration of large masses and have been notoriously hard to detect. That was until the LIGO Michelson interferometer in Hanford and Livingston was complete in 2015. A Michelson interferometer is a device that uses the interference of two beams of light to detect small changes in the path distance of the two beams. A diagram of one can be seen in Figure 1. By

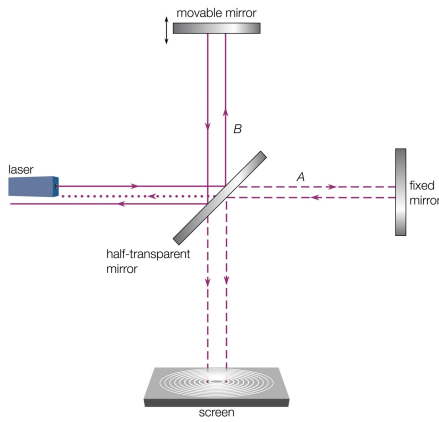


Figure 1: Diagram of a michelson interferometer as used in LIGO.¹

using a Michelson interferometer in the LIGO experiment the small changes in distance that are required can be detected and measured. These distances can be on the order of 10^{-21} m this distance is called the strain of the wave and is the amount of stretching over the original length and can be approximated as

$$h \approx \frac{GM}{c^2 d} \left(\frac{v}{c} \right)^2 \quad (1)$$

where G is the gravitational constant, M is the mass of the source, c is the speed of light, d is the distance to the source and v is the velocity of the system.

caused by the passing of gravitational waves

moving at the speed of light through the interferometer arms (which results in a shift in the interference pattern of the light beams). The first detection of a gravitational wave was on the 14th of September 2015, just 100 years after the publication of Einsteins paper. The first detection was of a binary black hole merger, these mergers commonly release a large amount of energy in the form of gravitational waves. This happens because as the two black holes accelerate towards each other they warp the space-time around them, and as they approach the point of coalescence the amplitude of these waves massively increases, thus allowing them to be detected over the Background noise. The run-down after merging is extremely quick and thus leaving a distinct peak at the time of coalescence. In this report the first 11 detections of gravitational waves as a result of binary black hole mergers will be analysed and discussed. Starting with GW150914.

GW150914

To be able to carry out the analysis of the gravitational waves it was necessary to set up our workspace to be able to use some provided function and packages written by the LIGO collaboration. This was done by first installing the LIGO lalsuite package for Python 3.10 and then importing the packages into our notebook. The package contains a number of useful functions that will be discussed in more detail later in this report. For the first detection of gravitational waves GW150914 the data was provided by the university through the Jupyter Hub. Once this data was loaded in the first step was to plot the strain against the time and identify by eye the peak of the gravitational wave. This plot is shown in Figure 2.

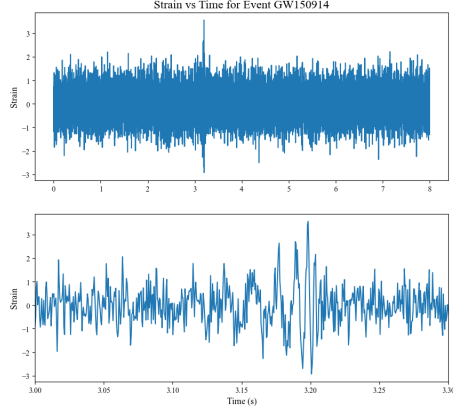


Figure 2: top: Plot of the strain against time for GW150914 bottom: limited to shorter time to resolve the peak more.

From Figure 2 it can be seen that the peak of the gravitational wave occurs at around 3.2 seconds. Knowing this time we can use the SCIPY package to generate a spectrogram of strain and frequency and from this a color plot can be created which visualizes the amount of energy in the gravitational wave at a given frequency and time. This plot is shown in Figure 3.

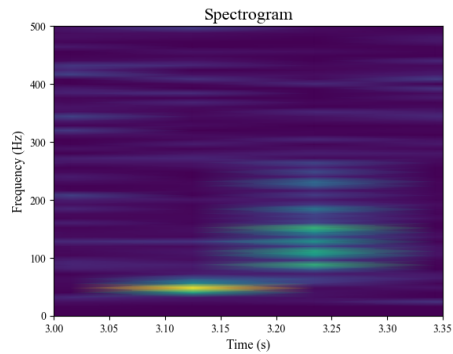


Figure 3: Spectrogram of GW150914, showing the 'chirp' track of the gravitational wave.

In this plot the point where the energy is high across multiple frequencies correlates with the same time as the peak in the strain in Figure 2.

Now that we have visualised the actual data from the gravitational wave, it would now be useful to be able to compare this to the theoretical prediction of the event. This can be generated using the make template function as supplied by the LIGO collaboration. This function takes in the masses of the two black holes and the time, frequency, distance and uncertainty on the data. The function then returns a strain and time array that can be used to plot the theoretical predictions. This produces an ideal signal as seen in Figure 4, this is easier to see the signal as, it no longer has any noise.

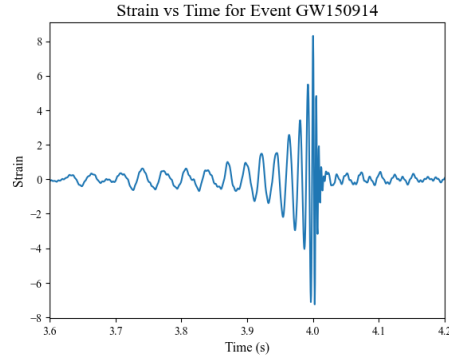


Figure 4: Theoretical prediction for GW150914

Later this template will be overlaid on the true data to determine the goodness of fit.

Detecting the signals for the ten Black hole Mergers

To determine the signal of the data for the ten black hole mergers the signal to noise ratio was used to determine the strength of the signal and its location in time. This could be done using the function given in the lalsuite package. This function returns a list of signal to noise ratios for each time in the data. If the SNR is low then the template that was put into the function is badly aligned or the masses of the black holes that were put into the template function are incorrect, and

if the SNR is high then the template is aligned well and the masses are correct. An example SNR time series is shown in Figure 5. This SNR

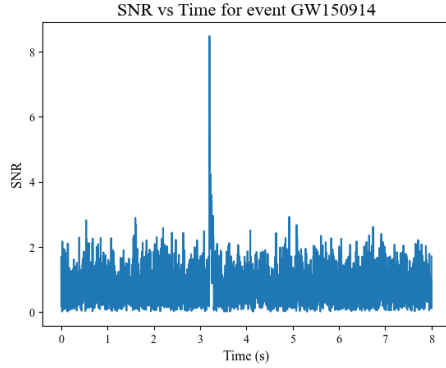


Figure 5: SNR time series for GW150914, with the peak at 3.2 seconds which tells us it is correctly aligned as this is the same time as obtained above

at the peak is important as it can be used later to determine the masses of the black holes as well as other important information. So this value was stored to be used later.

Determining the masses of the black holes

To determine the masses of the black hole pairs, some assumptions were made about the system. The first assumption was that m_1 the mass of the first black hole is larger than m_2 the mass of the second black hole. And the second assumption was that m_1/m_2 is less than 8 as this is the maximum mass ratio expected and it helps lower the number of possible mass pairs for the next section. With these assumptions the mass of the black holes can be determined using nested for loops to iterate over 10,000 possible mass pairs and calculate the SNR for each pair, then storing this SNR and the corresponding mass pair. This was done for the 10 binary black hole mergers (the binary neutron star merger was not included

in this section as it has different mass ranges, and thus will be analysed separately later). With the SNR's and the mass pair options stored a color plot was created for each event, where the intensity of the color is proportional to the SNR and the x and y axis are the mass of each black hole. An example of one of these plots is shown in Figure 6.

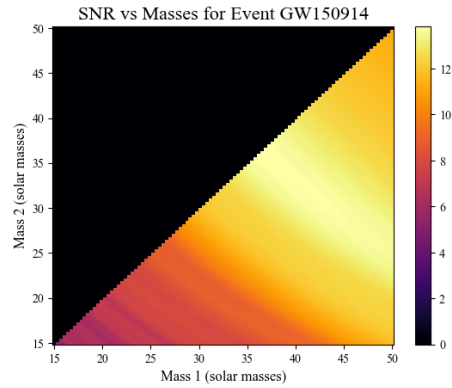


Figure 6: Mass plot for GW150914, showing the mass pairs that give the highest SNR

The black upper triangle is due to the conditional statement that mass one must be greater than mass two. Where the color is the highest is where the SNR is highest and thus the mass pair that is most likely to be correct. From this plot and as discussed above the mass pair that is most likely to be correct is 36.06 ± 1.4 and 35.74 ± 4.3 solar masses. This is within estimated error of to the masses published by the LIGO collaboration of 35.6 and 30.6 solar masses

Determining estimates for the distance, time and phase of coalescence

To find estimates for the distance, time and phase of coalescence the make template function was used again to generate a template for each event but this time using the masses as found above, as

Event	Distance (Mpc)	Time (s)	Phase (rad)	Mass one (M_{\odot})	Mass two (M_{\odot})
GW150914	1128 ± 82	3.198 ± 0.000	0.03 ± 0.16		
GW151012	3783 ± 1677	2.703 ± 0.001	0.00 ± 0.82		
GW151226	1812 ± 552	2.067 ± 0.000	0.00 ± 0.46		
GW170104	2119 ± 340	1.975 ± 0.000	1.79 ± 0.32		
GW170608	664 ± 81	3.794 ± 0.000	2.67 ± 0.18		
GW170729	3150 ± 590	3.539 ± 0.001	2.89 ± 0.51		
GW170809	2146 ± 444	2.641 ± 0.001	2.78 ± 0.42		
GW170814	2268 ± 385	3.335 ± 0.000	0.00 ± 0.34		
GW170818	1596 ± 417	2.161 ± 0.000	0.73 ± 0.40		
GW170823	4111 ± 1121	4.456 ± 0.001	0.00 ± 0.64		

Table 1: Estimated values for the distance, time and phase of coalescence for the ten binary black hole mergers

this would produce the most accurate templates for the sets of data. From this the SCIPY curve fit function could be used as this allow for estimated of unknown quantities depending on the how close the tru data/plot is to the template. However for this particlar application it was necessary to provide some upper and lower bounds that these value could take as this reduces the possible sollutions to the problem. These bounds were determined by looking at the peak and its location as well as an overall generic distance range for the events. The curve fit function also provides a covariance matrix which can be used to determine the error on the estimated values. This was repeated for each event and the results are shown in Table 1.

Analysis

Conclusion

References