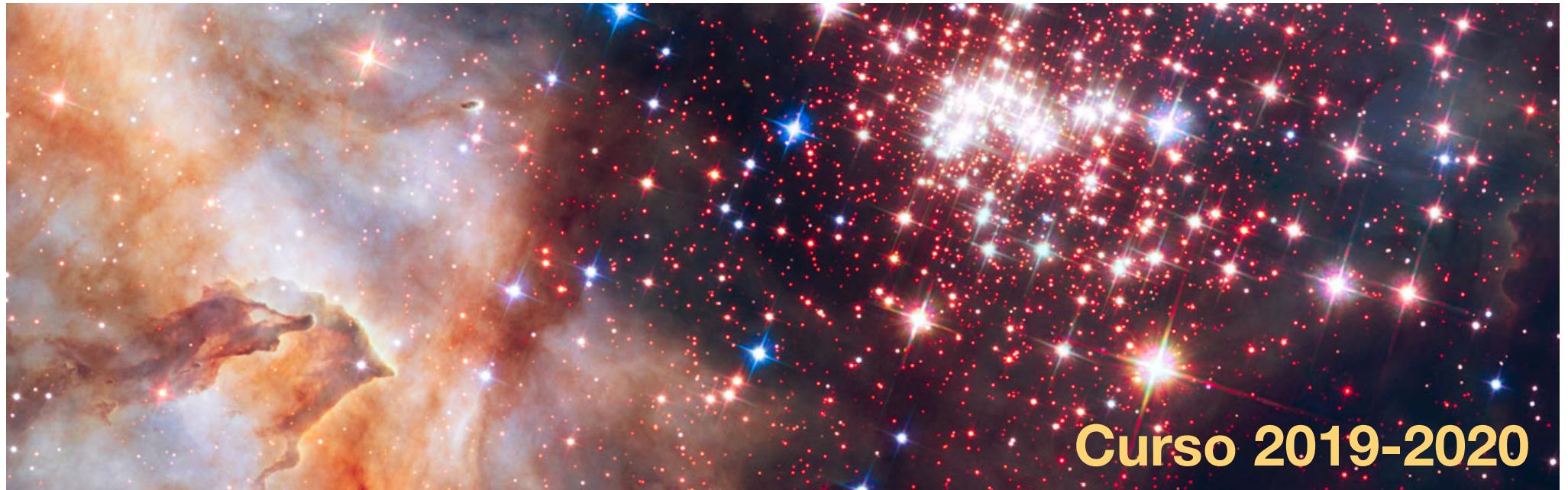


Física Estelar

Lluís Galbany, Ed. Mecenas (#16)

Inma Domínguez, Ed. Mecenas (#17)

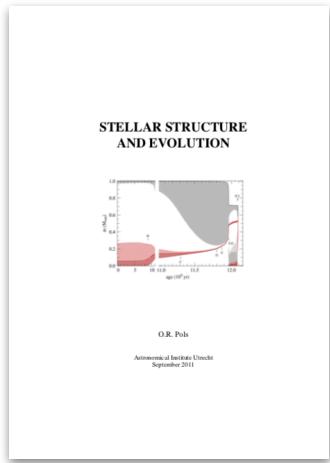
Antonio García, Ed. Mecenas (#16)



Curso 2019-2020

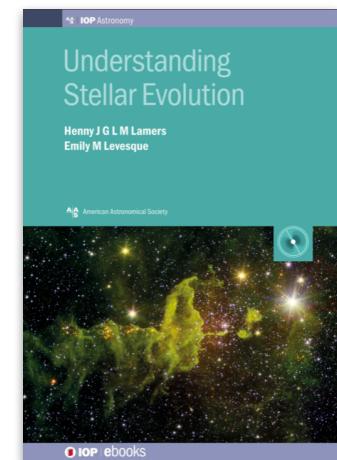
Syllabus

Aim: to understand the structure and evolution of stars, and their observational properties, using known laws of physics.



Stellar structure and evolution
(O. R. Pols)

Understanding Stellar Evolution
(Henny J. G. L. M. Lamers & Emily M. Levesque)



Goals

In this course we will:

- understand the global properties of stars: energetics and timescales
- study the micro-physics relevant for stars: the equation of state, nuclear reactions, energy transport and opacity
- derive the equations necessary to model the internal structure of stars
- examine (quantitatively) the properties of simplified stellar models
- survey (mostly qualitatively) how stars of different masses evolve, and the endpoints of stellar evolution (white dwarfs, neutron stars)
- discuss a few ongoing research areas in stellar evolution

Syllabus

Tema 1: Estructura Estelar

Principios de conservación y ecuaciones de estructura estelar. Condiciones de contorno y métodos de resolución. Ecuación de estado y opacidad. Transporte de energía: Radiación, convección y conducción. Criterios de estabilidad convectiva.

Tema 2: Fuentes de Energía Estelar

Teorema del Virial. Reacciones termonucleares y ritmos de reacción. Principales cadenas y ciclos de combustión nuclear. Otros procesos nucleares de interés astrofísico.

Tema 3: Evolución Estelar

Formación estelar y pre-secuencia principal. Límites de masa estelar: enanas marrones y planetas. Edad cero y secuencia principal. Estimación de edades de cúmulos estelares. Evolución en la rama de las gigantes: estrellas RGB y AGB. Formación de enanas blancas. Estrellas masivas y supernovas de colapso gravitatorio.

Tema 4: Objetos compactos

Evolución de enanas blancas. Ecuación de Volkov-Openheimer. Estrellas de neutrones: ecuación de estado. Agujeros negros. Aplicaciones.

Tema 5: Evolución Estelar en Sistemas Binarios.

Supernovas termonucleares: Aplicaciones cosmológicas. Binarias cataclísmicas. Novas. Erupciones de rayos X. Estrellas gigantes binarias y anomalías químicas.

Tema 6: Pulsaciones Estelares.

Pulsaciones esféricas adiabáticas y no adiabáticas. Mecanismos de pulsación. La banda de inestabilidad en el diagrama HR. Oscilaciones no radiales. Astroismología

1. Introduction

What is a star?

A star is...

[fill in student answers from discussion]

What is a star?

A star is an object that:

- (1) radiates energy from an internal source; and
- (2) is bound by its own gravity.

Stars must evolve: have *life* and *death*.



Stars can have only a limited range of masses: $0.1 \lesssim M_{\odot} \lesssim 100$



Fundamental properties of stars

The **Sun**, our nearest neighbor

Standard parameters of the Sun:

$$\boxed{\begin{aligned}M_{\odot} &\sim 1.99 \times 10^{33} \text{ g} \\R_{\odot} &\sim 6.96 \times 10^{10} \text{ cm} \\L_{\odot} &\sim 3.84 \times 10^{33} \text{ ergs}^{-1}\end{aligned}}$$

$$L = 4\pi R^2 \sigma T_{\text{eff}}^4 \longrightarrow T_{\odot, \text{eff}} \sim 5777 \text{ K}$$

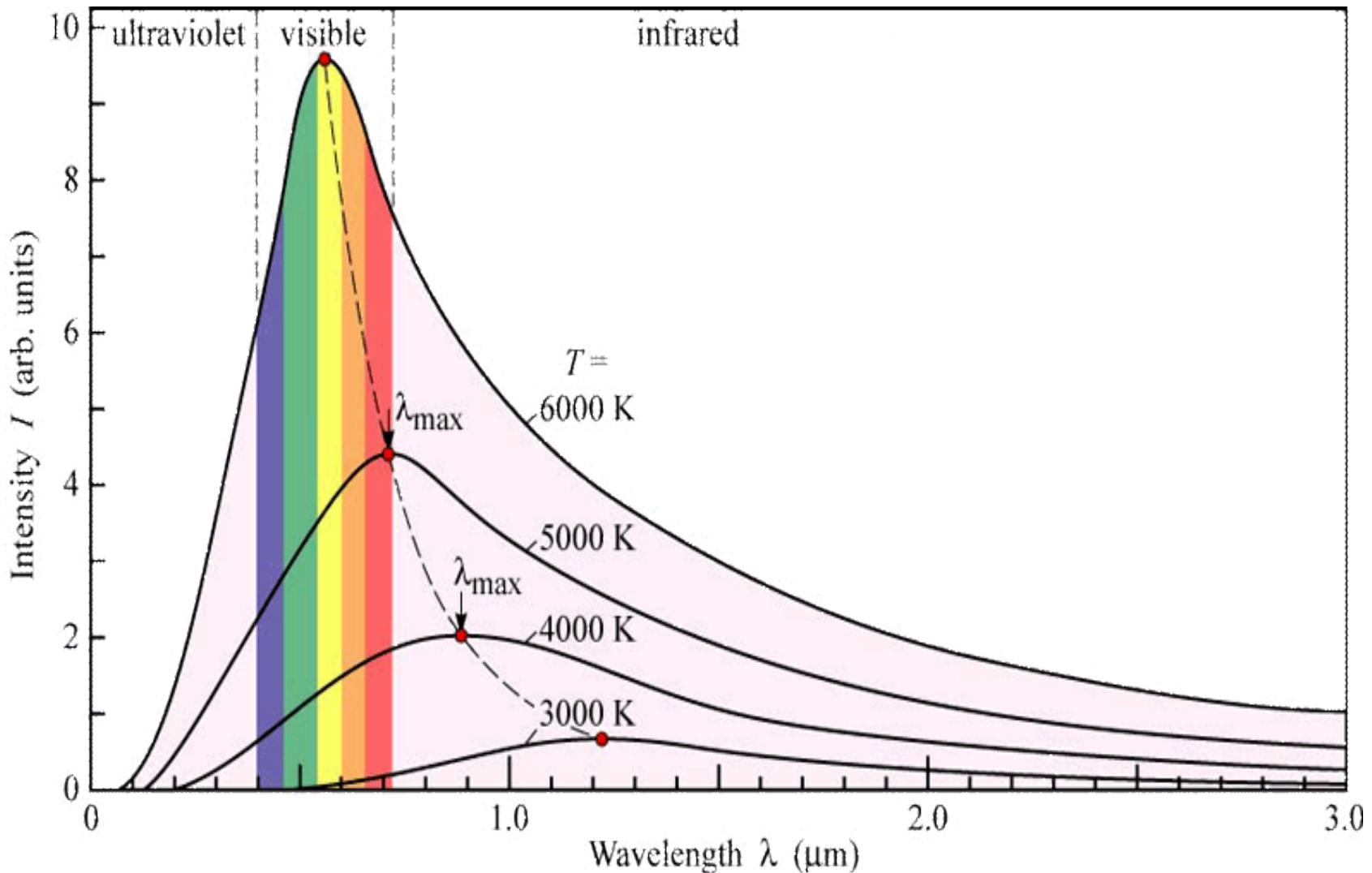
Solar **abundances** ($X + Y + Z = 1$) *(...also v_{rotation})*

Nr	Element	Z	m (AMU)	Abund	Nr	Element	Z	m (AMU)	Abund
1	H	1	1.0079	0.00	6	N	7	14.007	-3.95
2	He	2	4.0026	-1.01	7	Mg	12	24.305	-4.42
3	O	8	15.999	-3.07	8	Si	14	28.086	-4.45
4	C	6	12.011	-3.44	9	S	16	32.066	-4.79
5	Ne	10	20.180	-3.91	10	Fe	26	55.847	-4.46

$$H : He : C + N + O + Ne : \text{rest} = 0.70 : 0.28 : 0.016 : 0.003$$

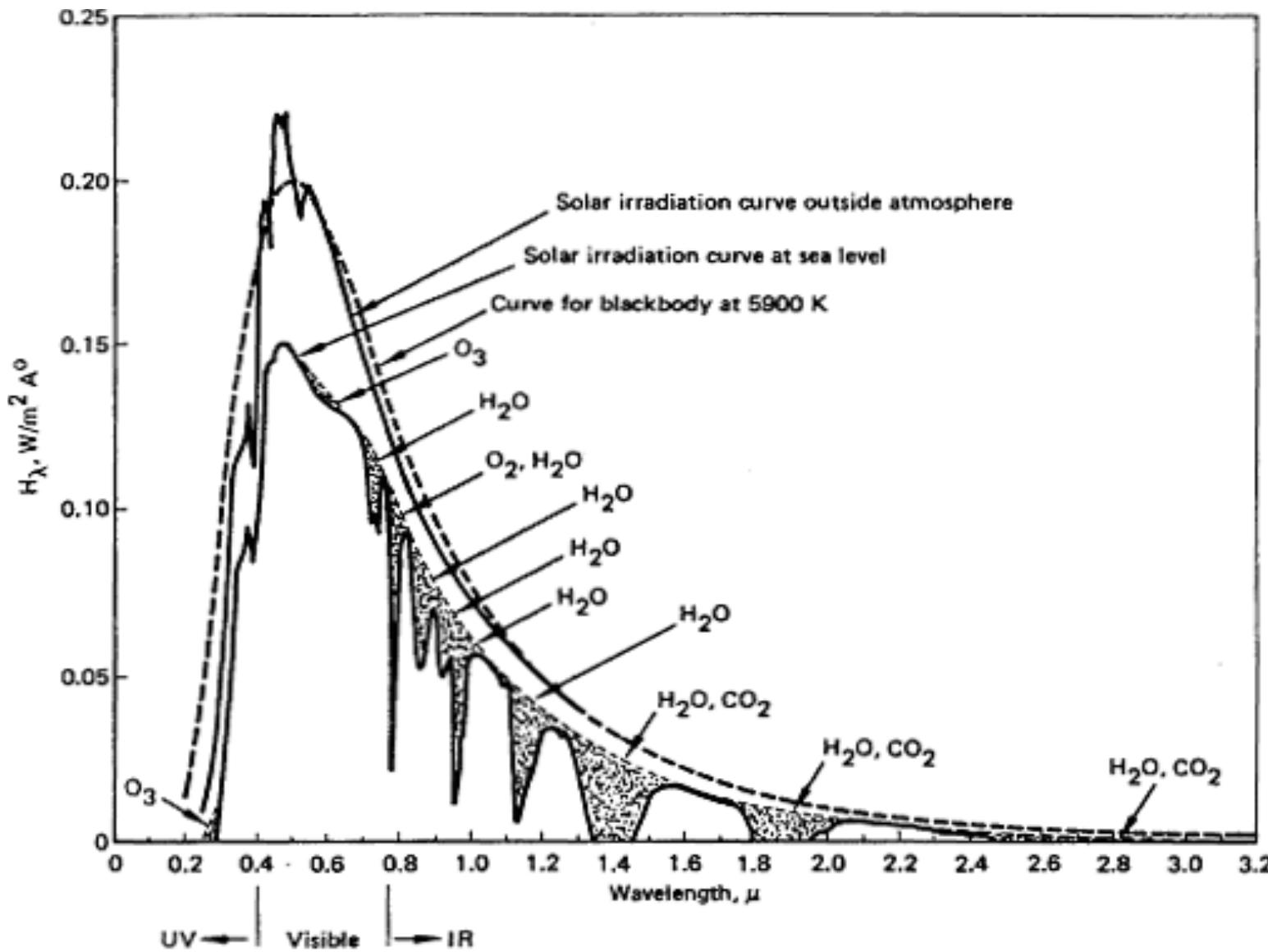
Effective temperature

Radiation curves of Blackbodies: $F = \sigma T^4 = \sigma(T_{\text{eff}}^4)$



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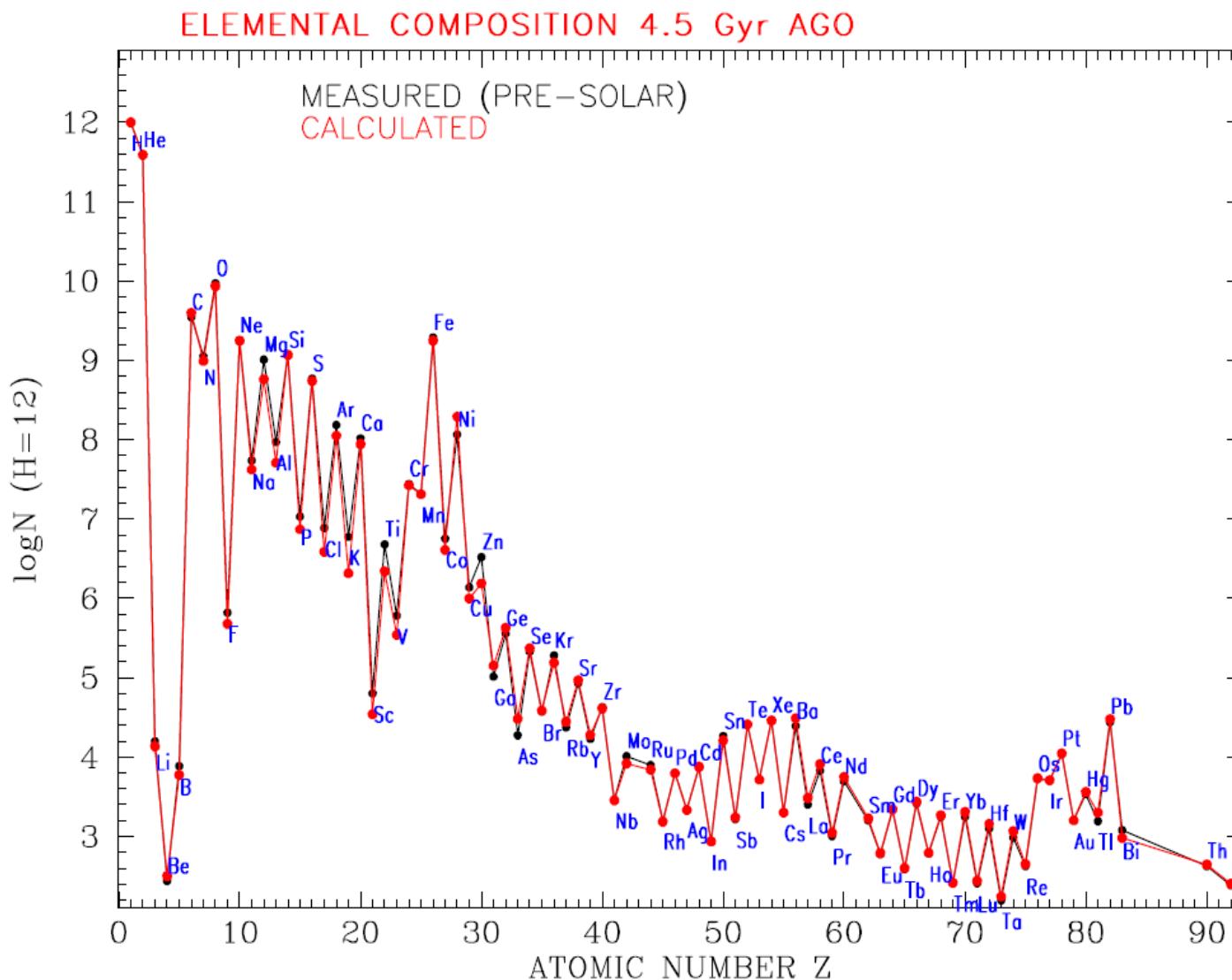
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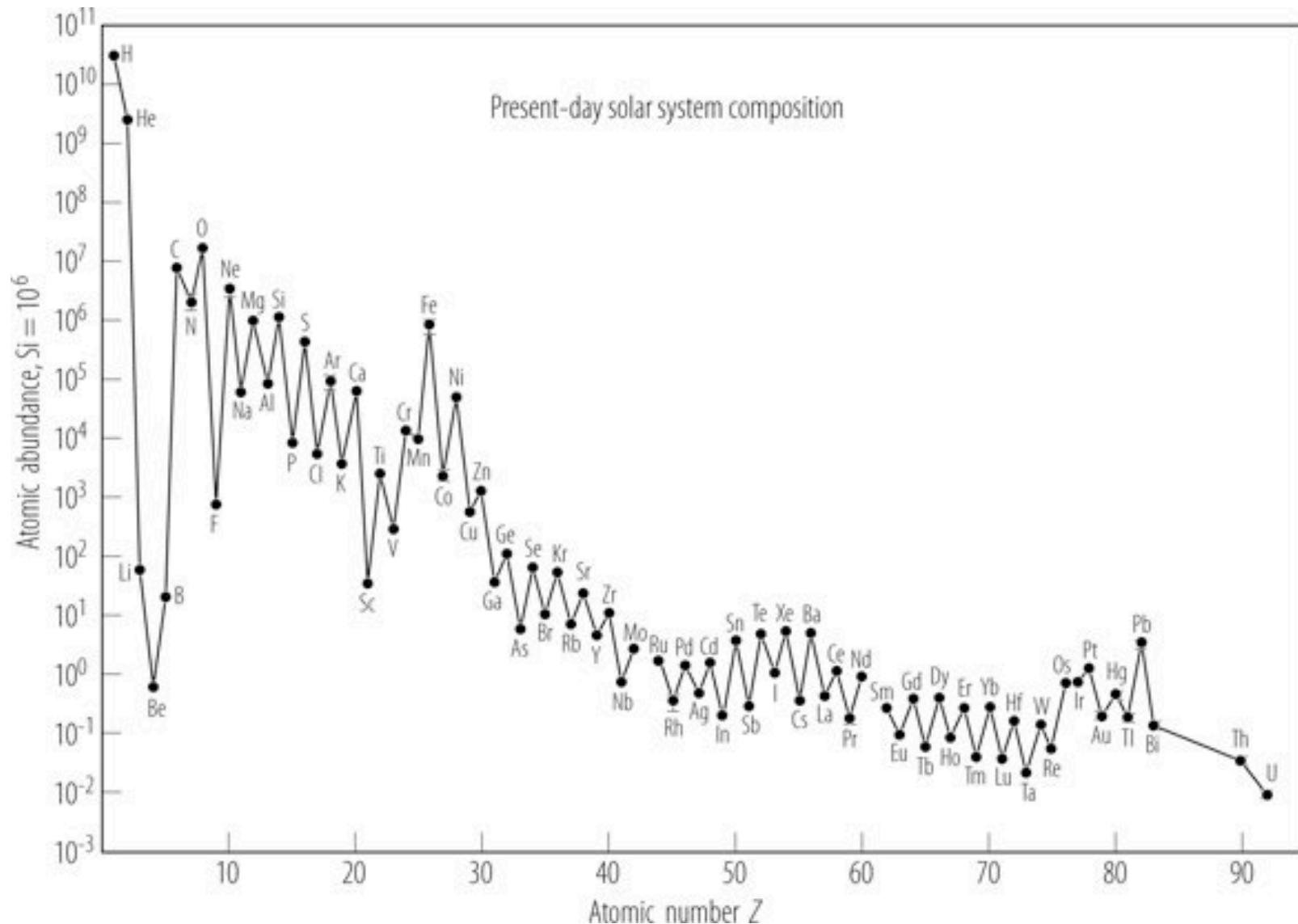
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Sun abundance



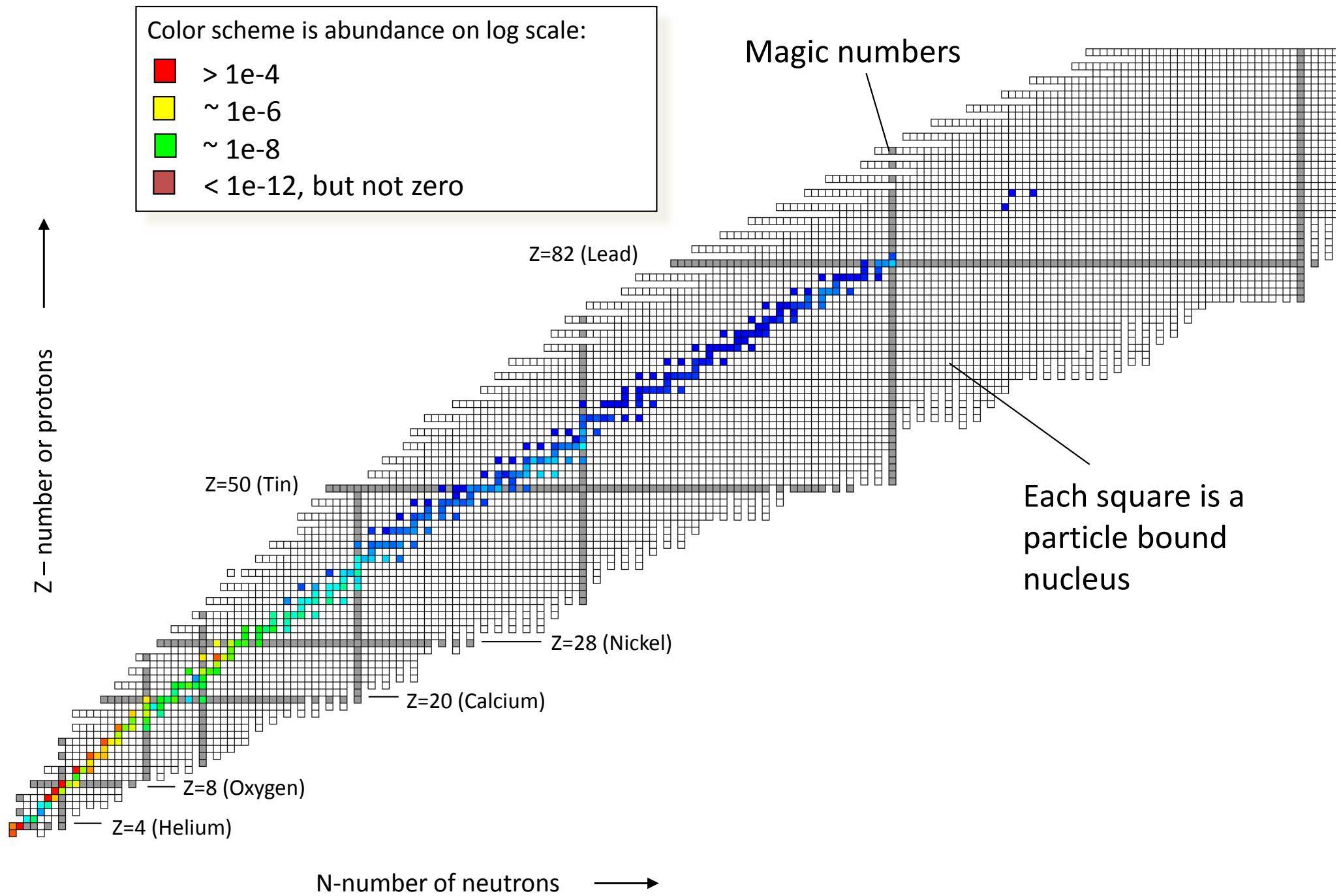
$A = \log(n_x/n_H) + 12$ (log of number of atoms per 10^{12} H atoms)
also used number of atoms per 10^6 Si atoms

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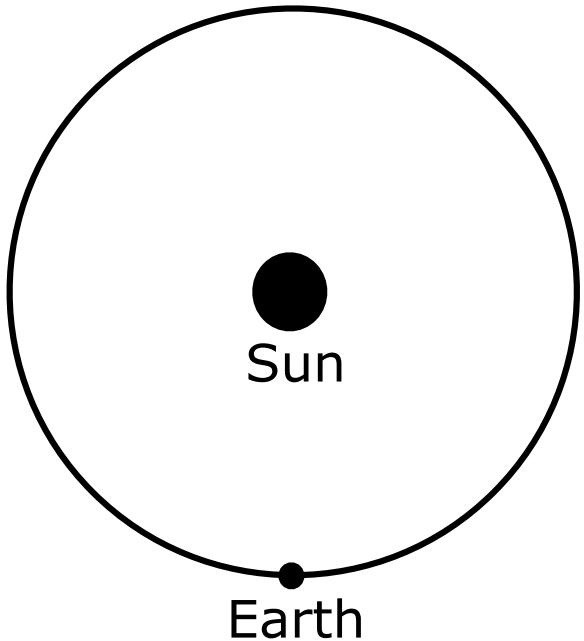
Abundances of nuclei on the chart of nuclides:



Observations of stellar parameters

Astronomical observations can yield information about stellar fundamental quantities:

1. Distance - parallax method

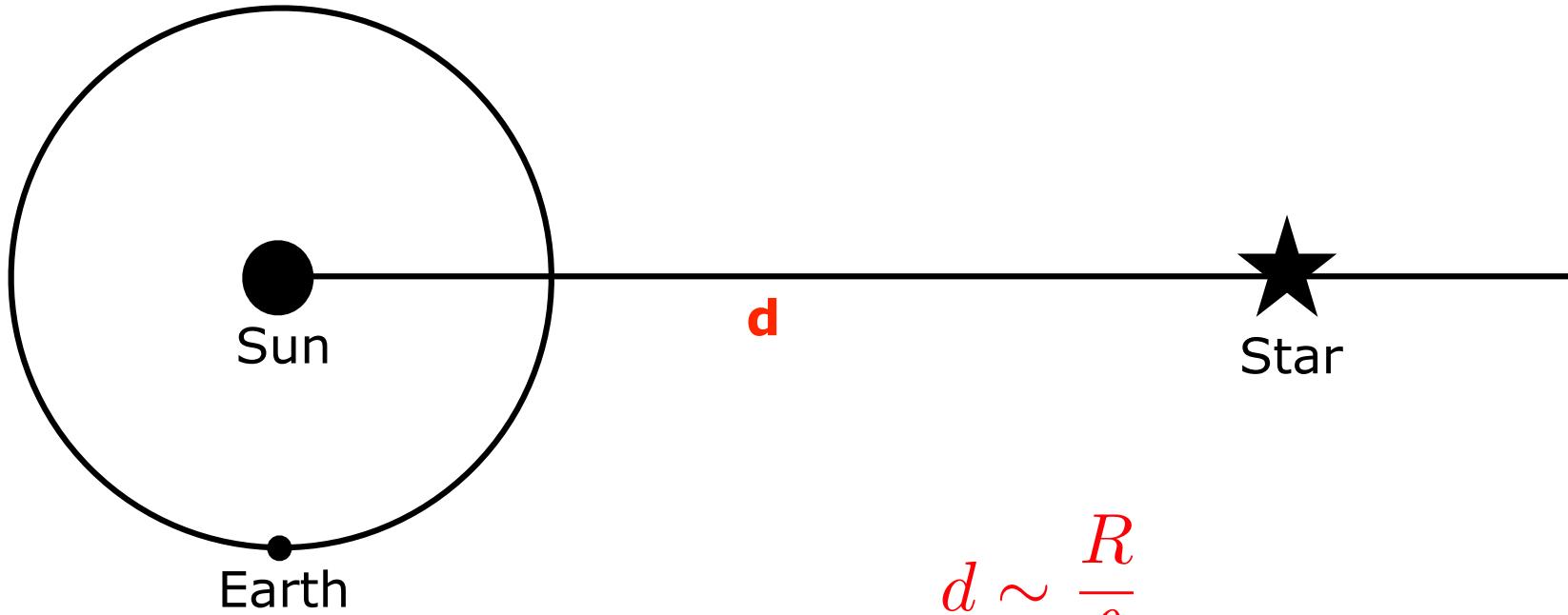


$$d \sim \frac{R}{\theta}$$

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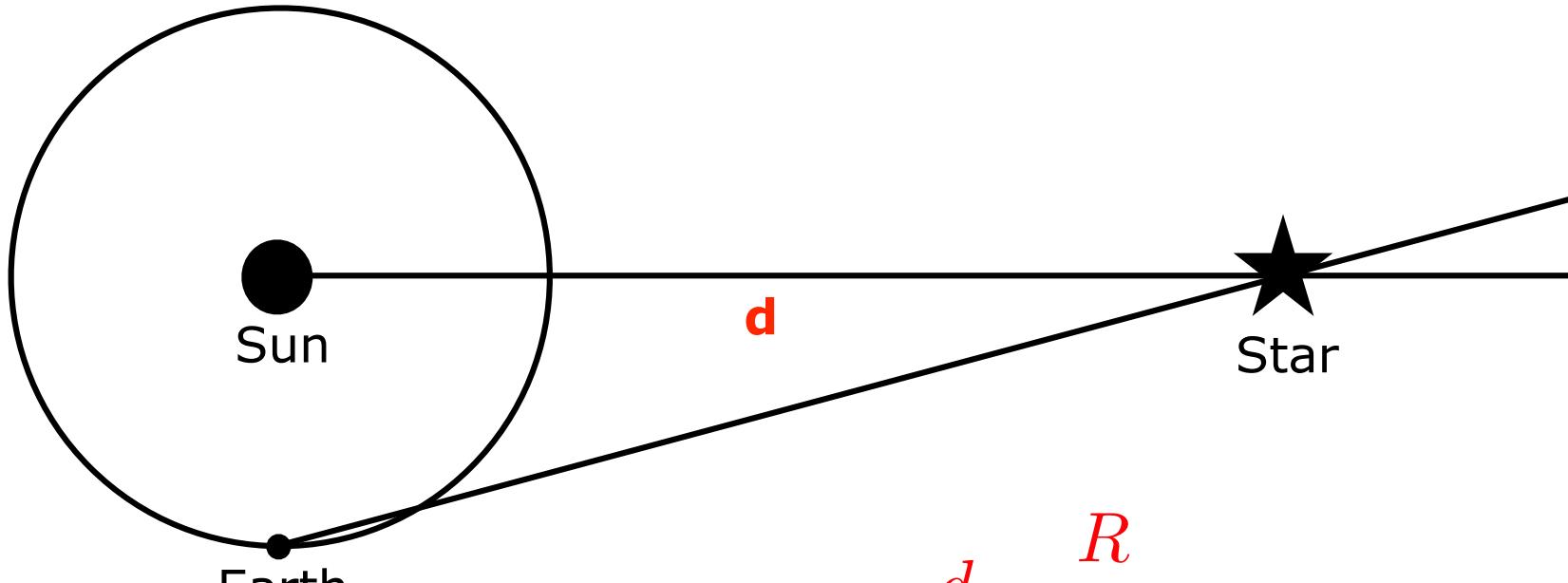
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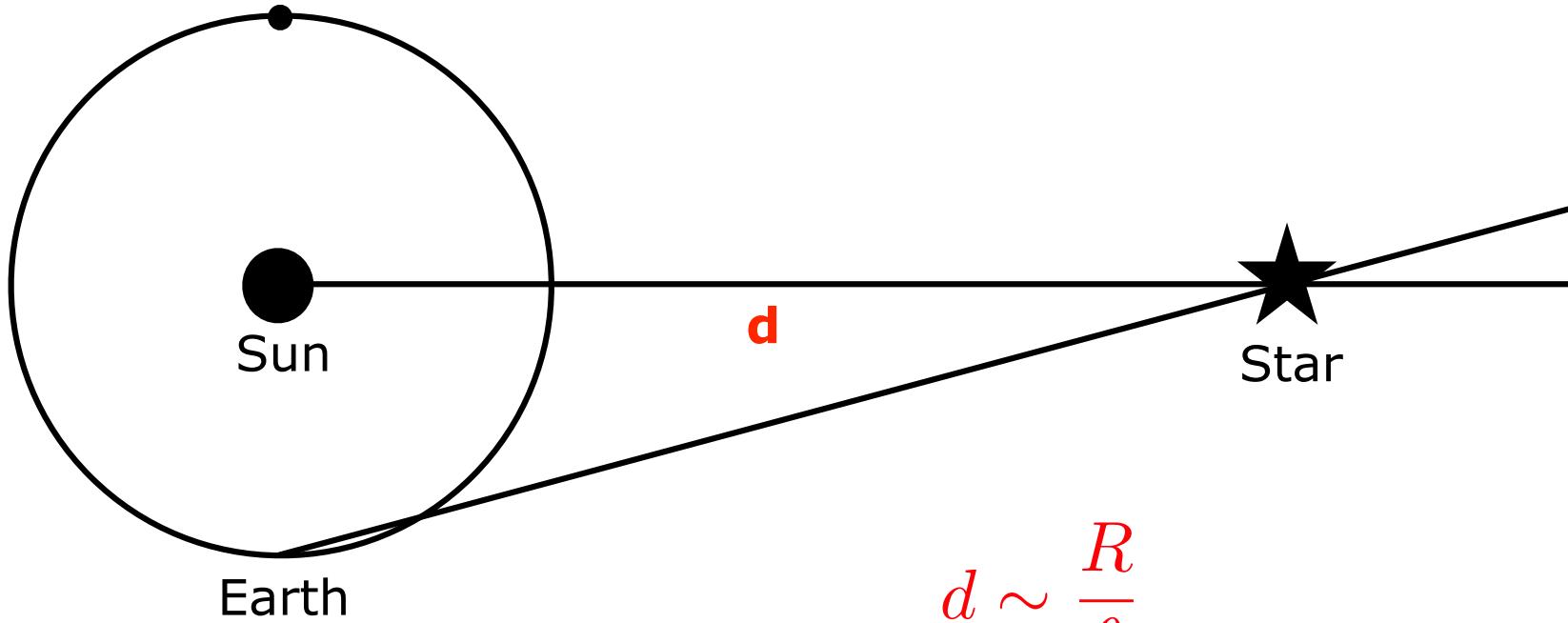
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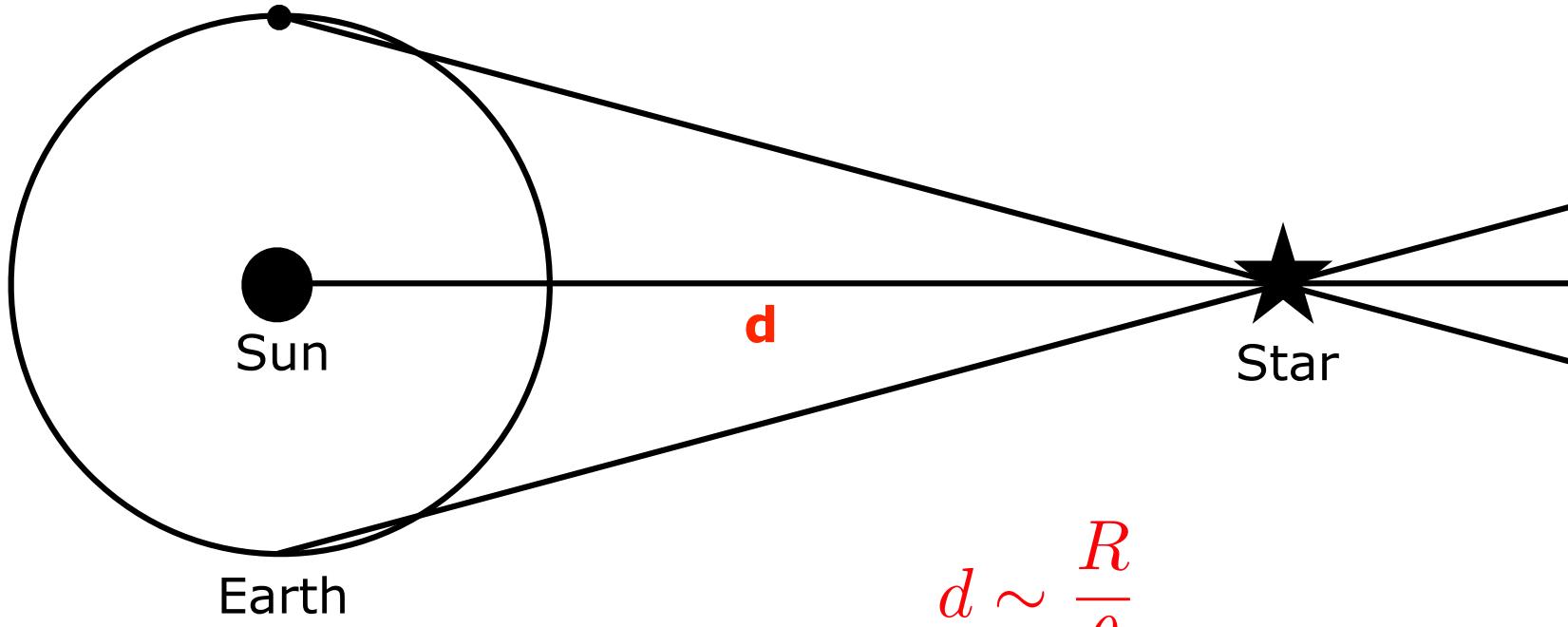
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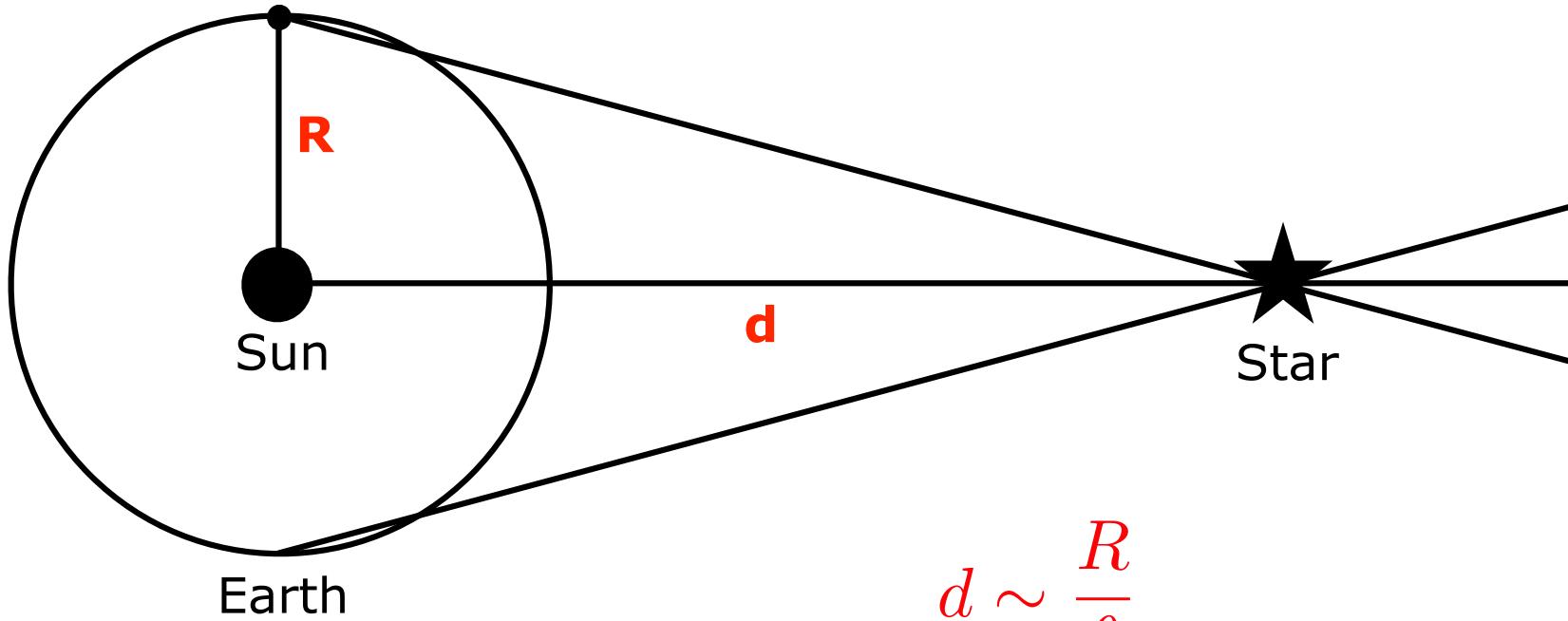
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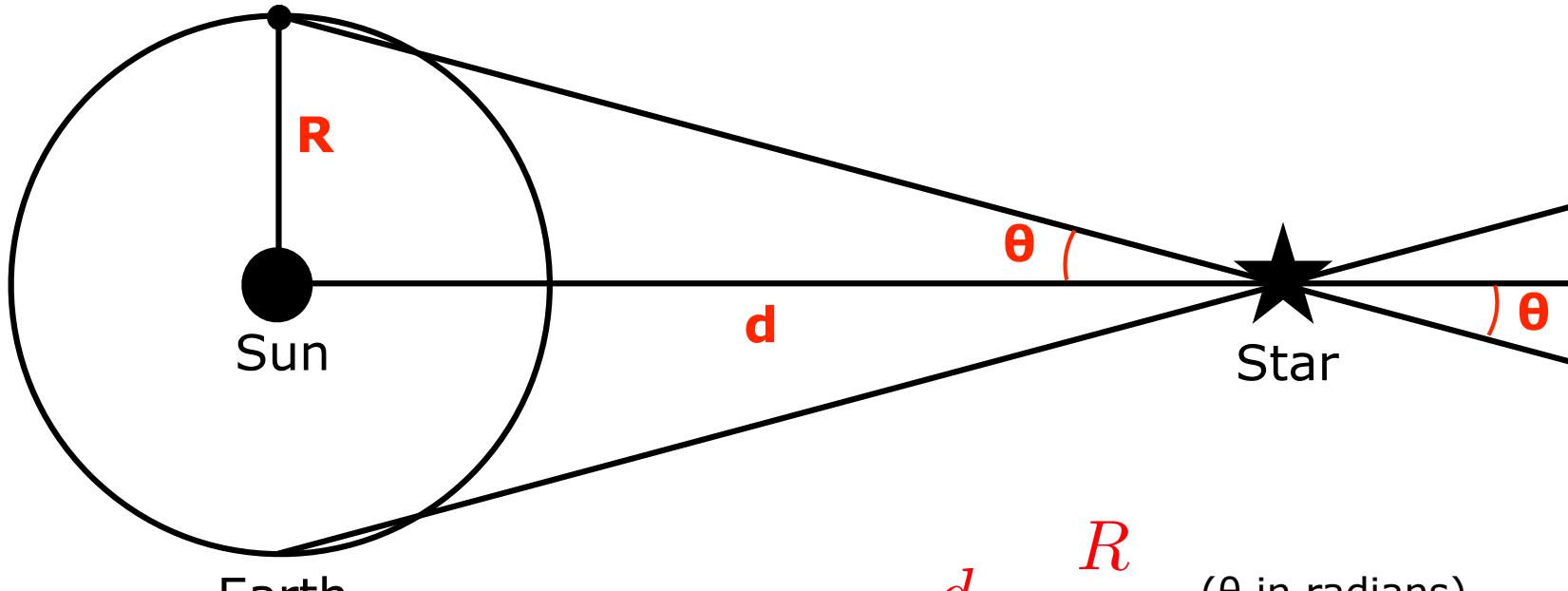
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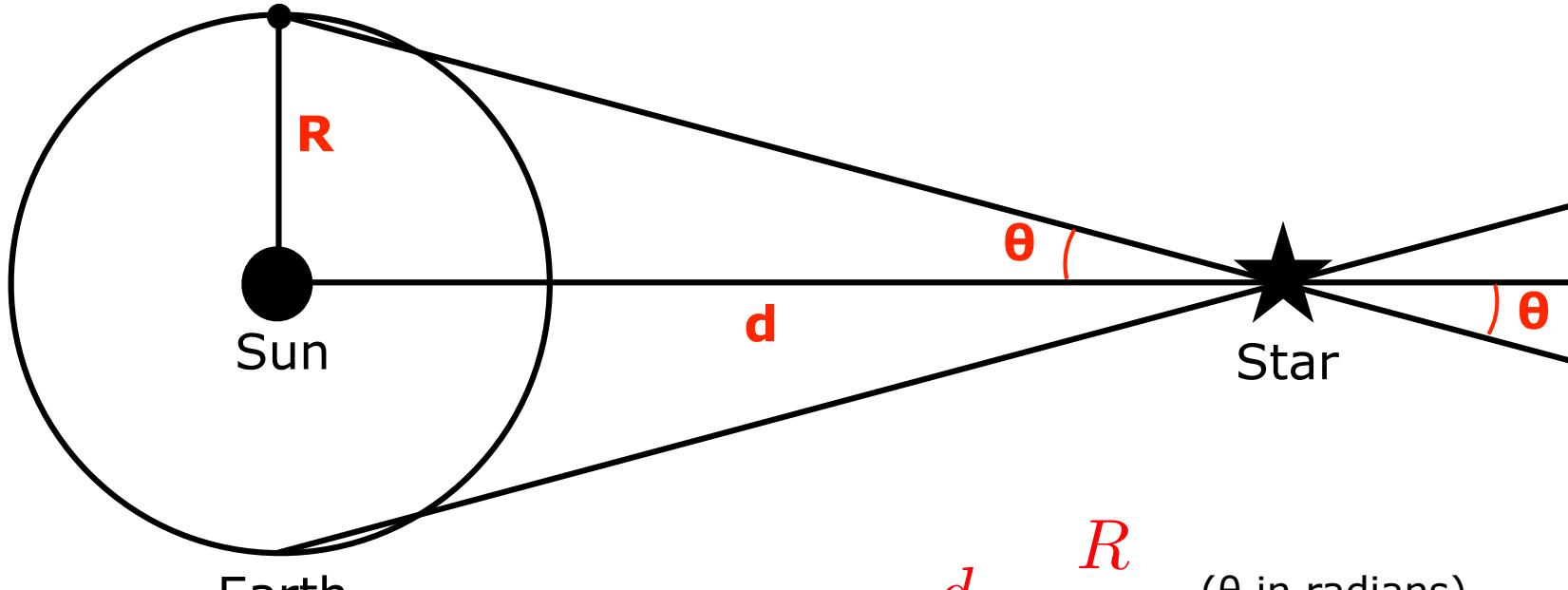
$$d \sim \frac{R}{\theta} \quad (\theta \text{ in radians})$$



Observations of stellar parameters

Astronomical observations can yield information about stellar fundamental quantities:

1. Distance - parallax method



$$d \sim \frac{R}{\theta} \quad (\theta \text{ in radians})$$



$$R = 1.5 \times 10^{13} \text{ cm}, 1'' = 4.85 \times 10^{-6} \text{ radians}$$

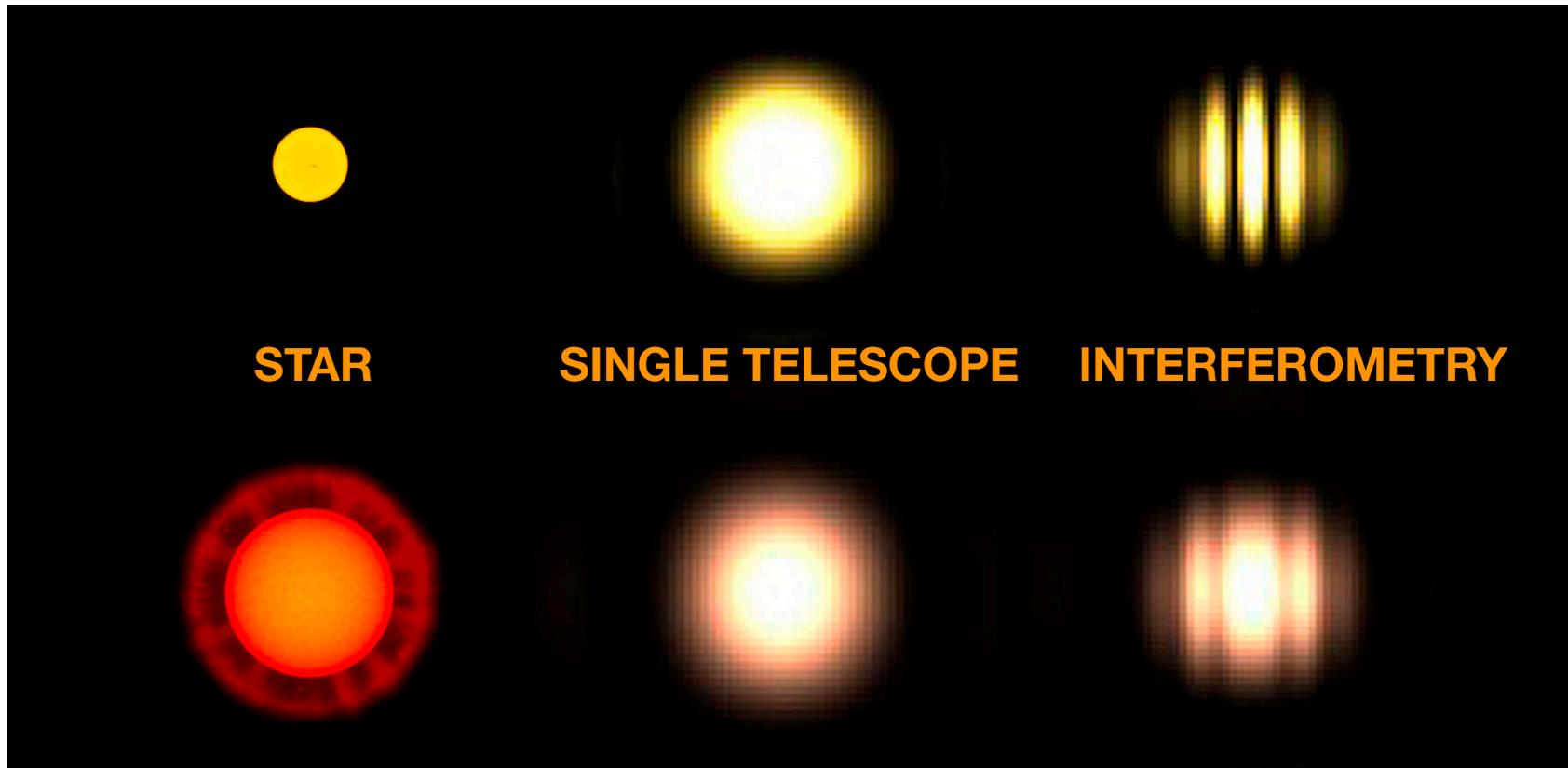
so a parallax of $1''$ = distance of $3.09 \times 10^{18} \text{ cm} = 1 \text{ parsec}$

Observations of stellar parameters

2. Angular radius

With interferometry, for stars sufficiently extended in the sky

$$\theta = \frac{R}{d} \quad \xrightarrow{L = 4\pi R^2 \sigma T_{\text{eff}}^4} \quad \sigma T_{\text{eff}}^4 = \frac{f_{\text{bol}}}{\theta^2}$$



Observations of stellar parameters

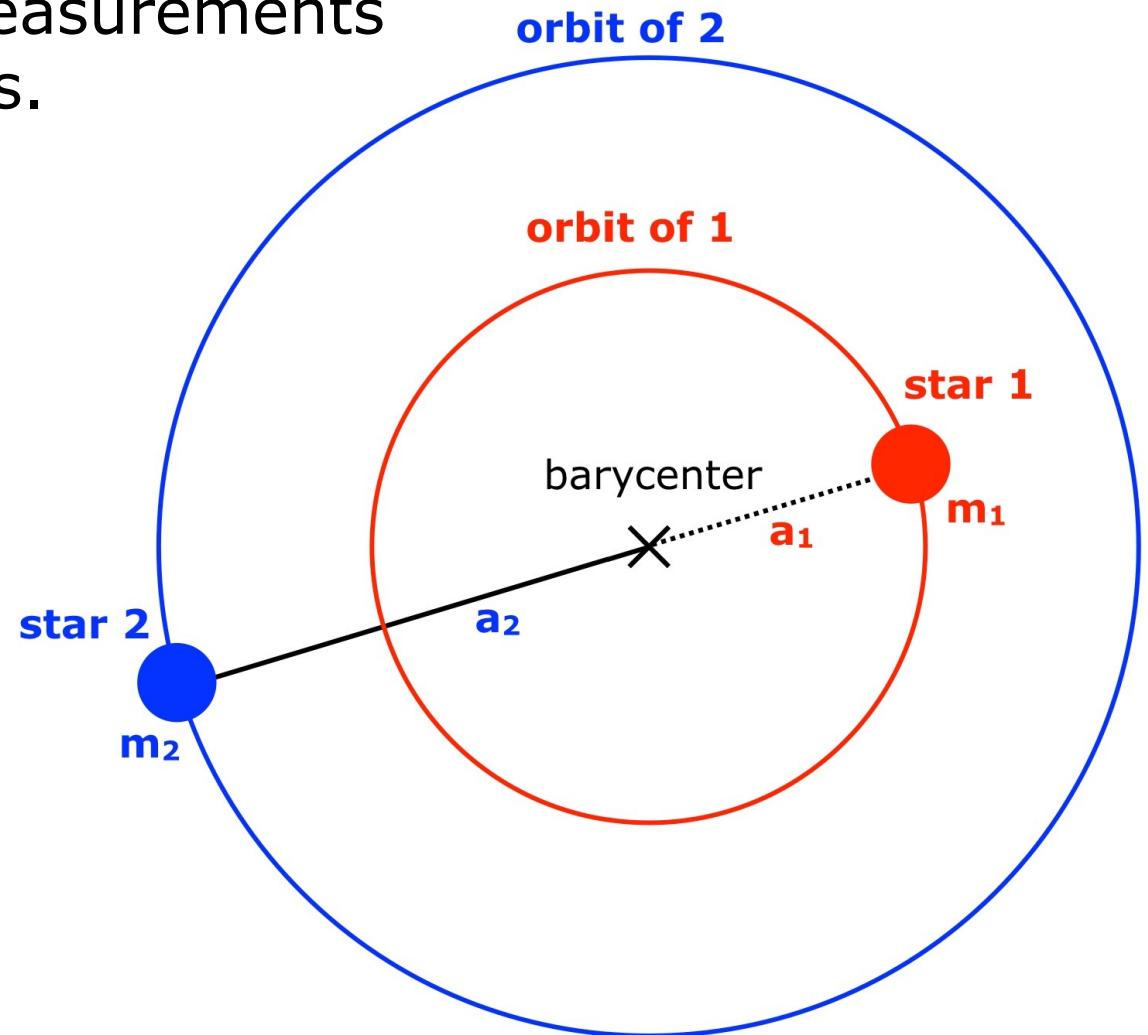
3. Mass

The most accurate mass measurements for stars come from binaries.

$$\left(\frac{2\pi}{P}\right)^2 = G \frac{(M_1 + M_2)}{a^3}$$
$$= G \frac{(M_1 + M_2)}{(a_1 + a_2)^3}$$

and

$$\frac{a_1}{a_2} = \frac{M_1}{M_2}$$



(requires orbits and radial velocities for both stars...)

Observations of stellar parameters

4. Luminosity

The luminosity (L) of a star can be derived from its distance and its radiative flux measured at Earth using the relations:

$$L = 4\pi d^2 f_{\text{bol}} = 4\pi R^2 \sigma T_{\text{eff}}^4 \quad \text{or} \quad \frac{L}{L_{\odot}} = \left(\frac{R}{R_{\odot}} \right)^2 \left(\frac{T_{\text{eff}}}{5777\text{K}} \right)^4$$

where

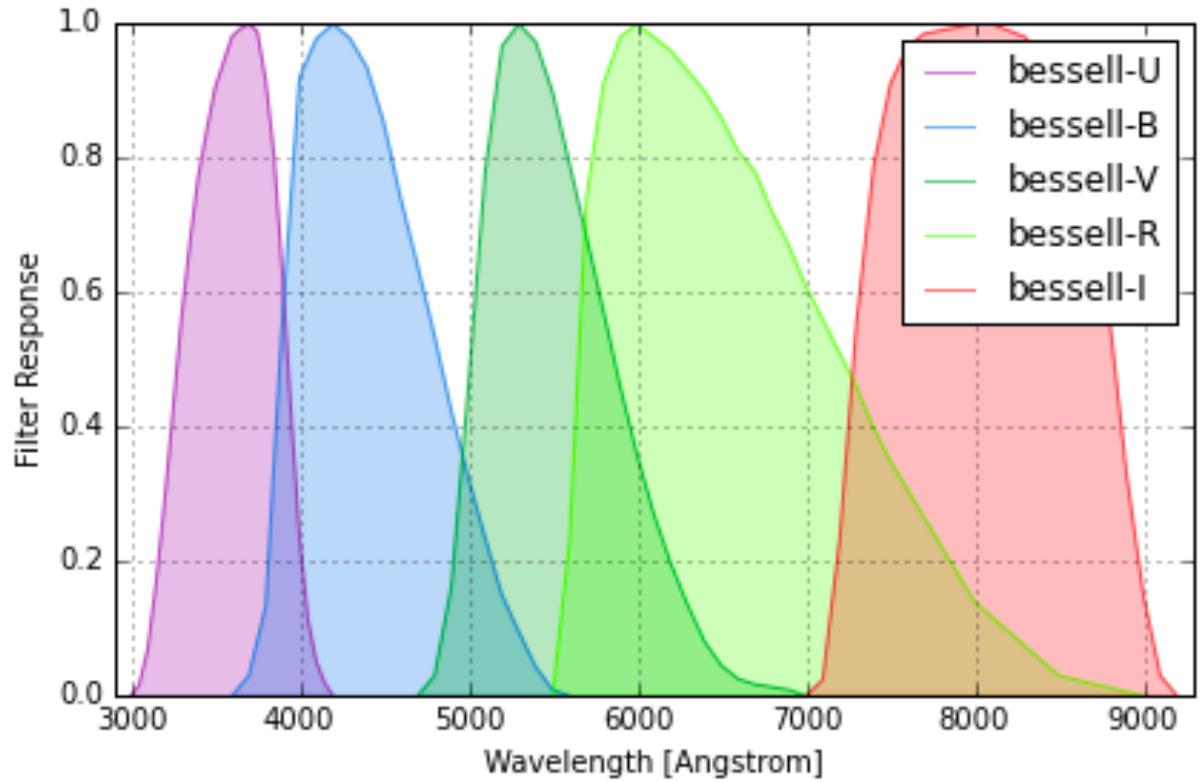
f_{bol} = flux integrated over all wavelengths and corrected for ISM and atmospheric attenuation

T_{eff} = approximation of the gas temperature in the photosphere of the star at optical depth $\tau \approx 1$

Observations of stellar parameters

5. Magnitude and color

Magnitudes reflect the energy flux received on Earth in different wavelength bands.



Color indices are defined as the ratio of the flux in two wavelength bands, and often expressed as the difference between two apparent magnitudes (m_λ) in different filters

Observations of stellar parameters

$$m_\lambda = -2.5 \times \log \left(\frac{f_\lambda}{f_{0,\lambda}} \right)$$

Note: a *lower* value of m_λ corresponds to a *larger* flux

where f_λ is in $\text{erg cm}^{-2} \text{ s}^{-1}$ and $f_{0,\lambda}$ is the reference flux at that λ

Magnitude systems (ref. fluxes) typically used are :

Vega, star

$$f_0(\text{U|B|V}) = 4.2|6.4|3.8 \times 10^{-8} \text{ erg cm}^{-2} \text{ s}^{-1} \text{ nm}^{-1}$$

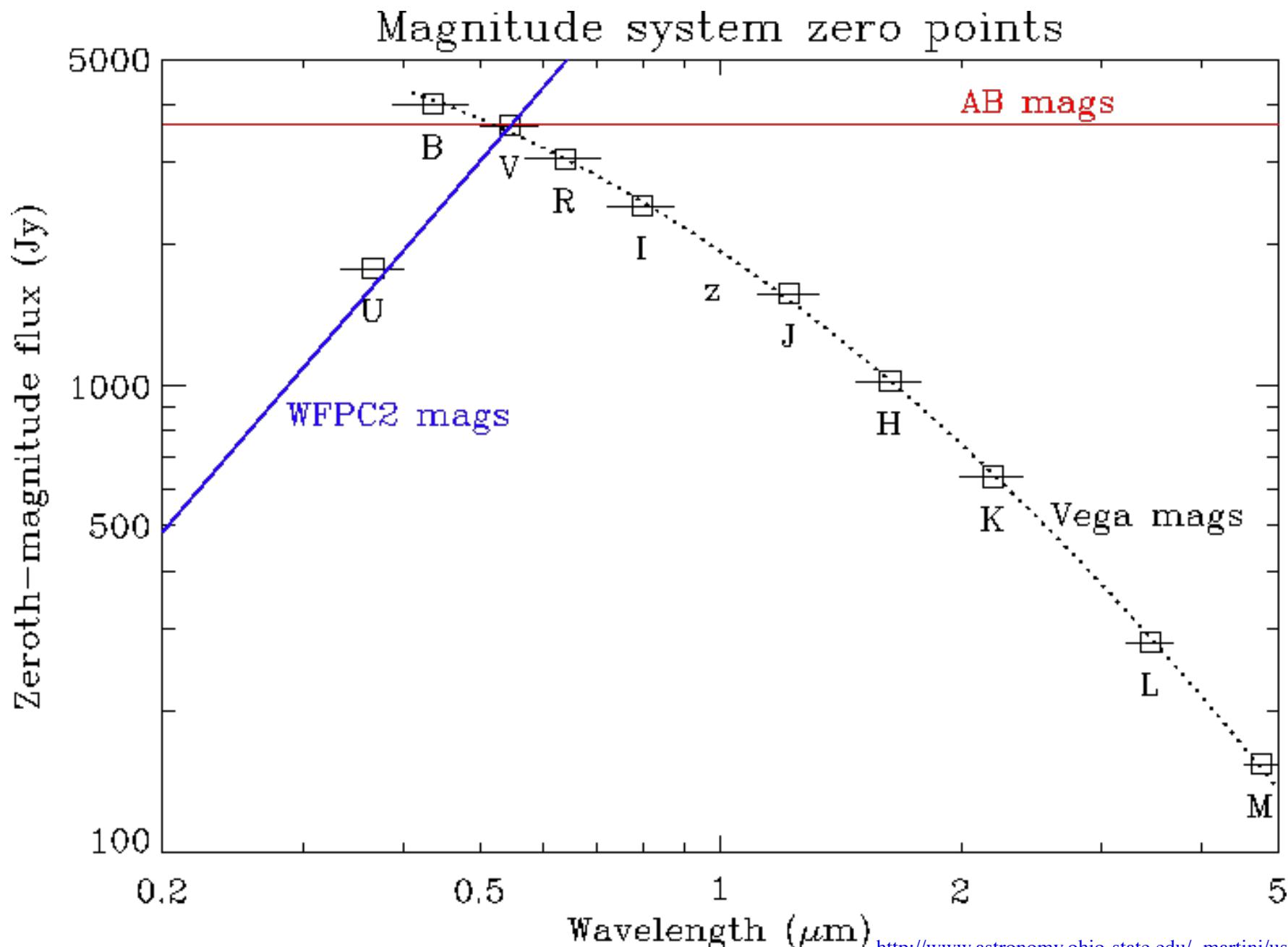
AB, flat $3.63 \times 10^{-20} \text{ erg cm}^{-2} \text{ s}^{-1} \text{ Hz}^{-1} = 3631 \text{ Jy}$

$$AB_{\text{mag}} = -2.5 \times \log F_\nu - 48.6$$

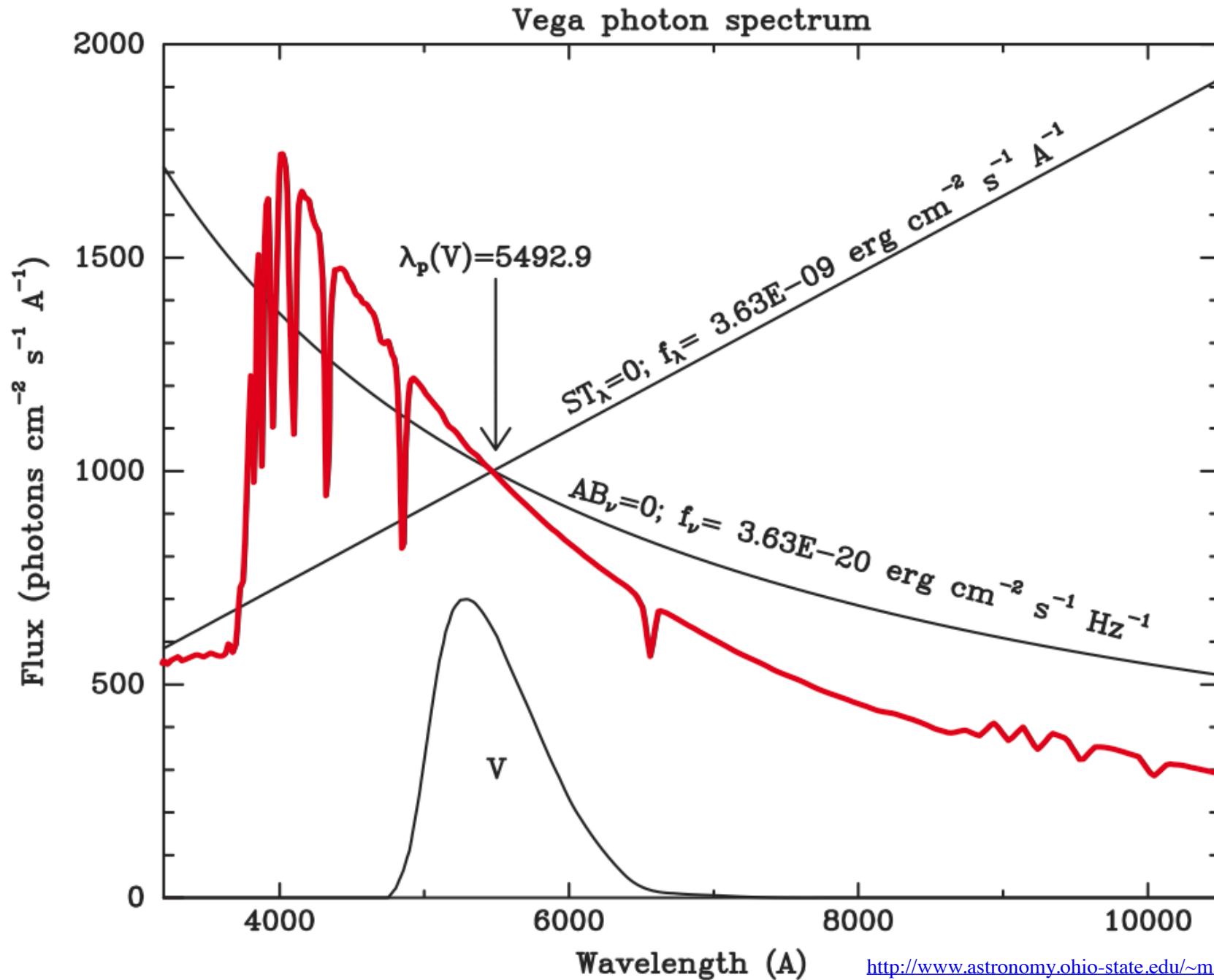
ST, flat $3.63 \times 10^{-9} \text{ erg cm}^{-2} \text{ s}^{-1} \text{ \AA}^{-1}$

$$ST_{\text{mag}} = -2.5 \times \log F_\lambda - 21.1$$

Observations of stellar parameters



Observations of stellar parameters



Observations of stellar parameters

Adopting the Johnson UBVRI filters, $m_V = V = -2.5 \times \log \left(\frac{f_V}{f_{0,V}} \right)$

B-V color index is $B - V = -2.5 \times \log \left(\frac{f_B}{f_V} \right)$ [+ difference in $f_{0,\lambda}$]
Ex. Calculate term
for Vega and AB.

Colors are defined as BLUE-RED mags. Following BB radiation, a **larger** (redder) value for color corresponds to a **lower** (cooler) T_{eff}

If a star's distance is known we can calculate its **absolute magnitude** M_λ (the m_λ the star would have at a distance of 10pc):

$$M_\lambda = m_\lambda + 5 - 5 \times \log(d)$$

The **absolute bolometric magnitude** M_{bol} describes the flux of the star integrated over all wavelengths at a distance of 10pc:

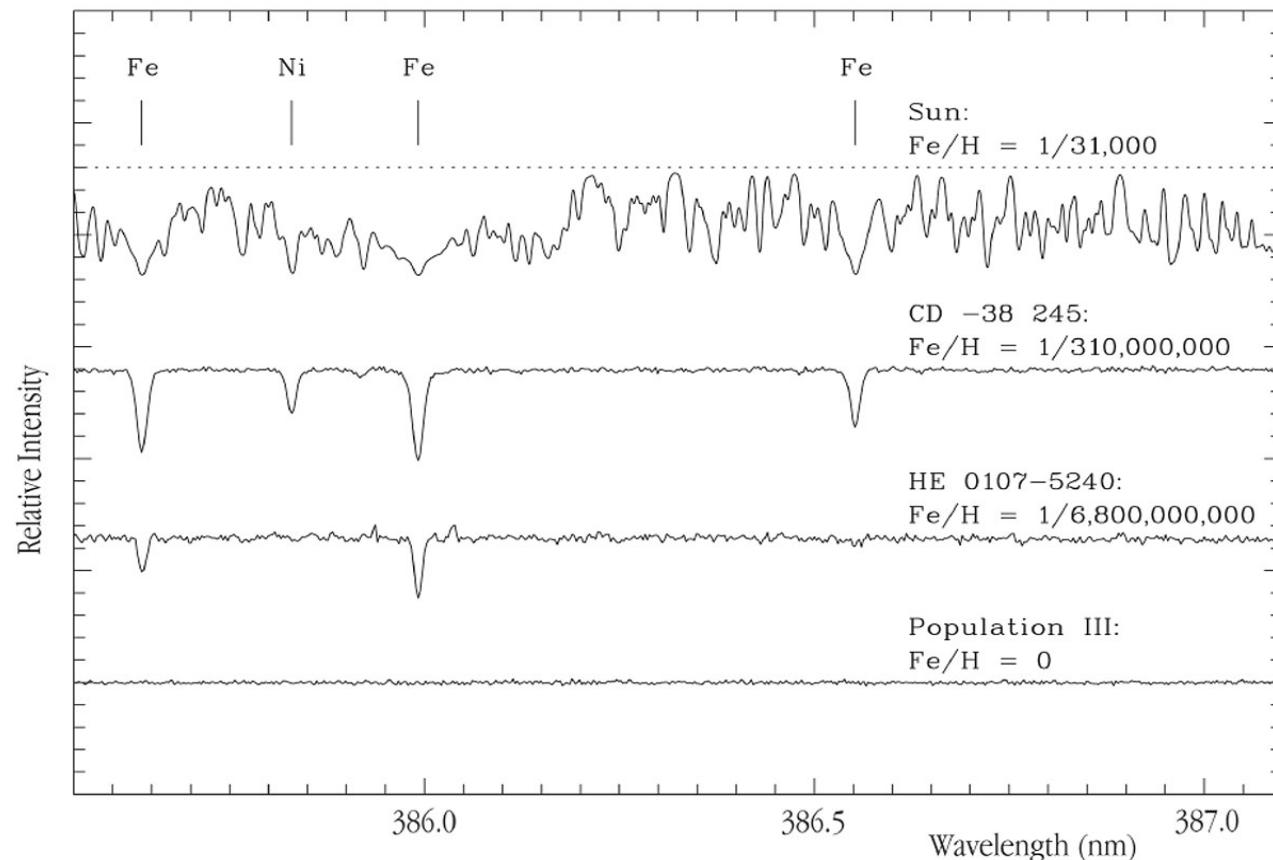
$$M_{\text{bol}} = -2.5 \times \log \left(\frac{L}{L_\odot} \right) + 4.74$$

and $M_{\text{bol}} - M_\lambda = BC_\lambda$, **the bolometric correction**

Observations of stellar parameters

6... Other spectral parameters

gravity, chemical abundances, Teff, rotational velocity...



Spectra of Stars with Different Metal Content

Observations of stellar parameters

So far, we listed *surface* stellar parameters

We need a theory of stellar structure to derive the internal properties of a star.

However, direct windows on the *interior* of a star exist:

Neutrinos that escape without interaction. Only detected from the Sun

Seismology, the frequency of different stellar oscillations contains information about the speed of sound waves inside the star, so do density and temperature.

The Structure of Stars

To get started...

Let's define some standard regions in a star...



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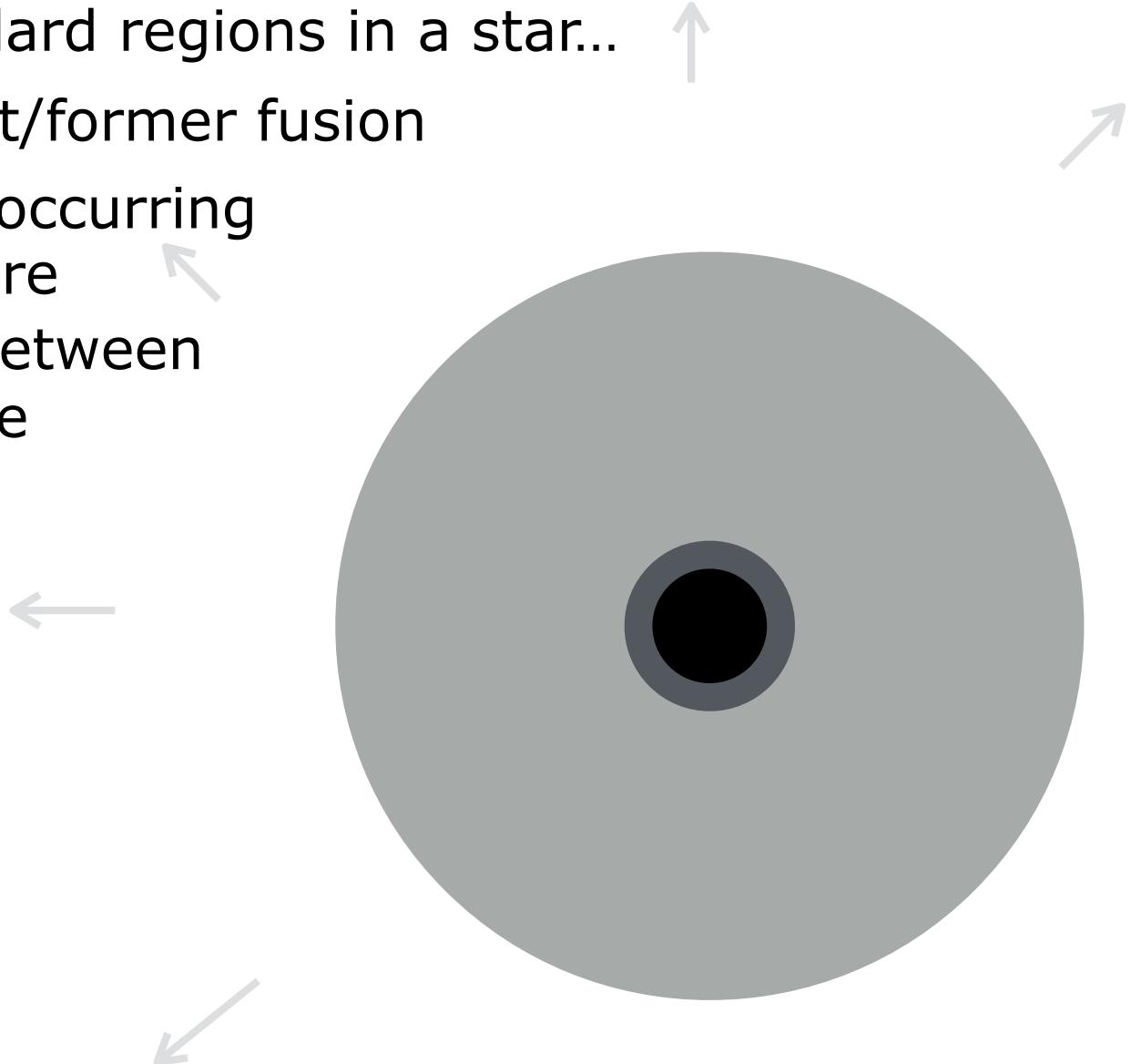


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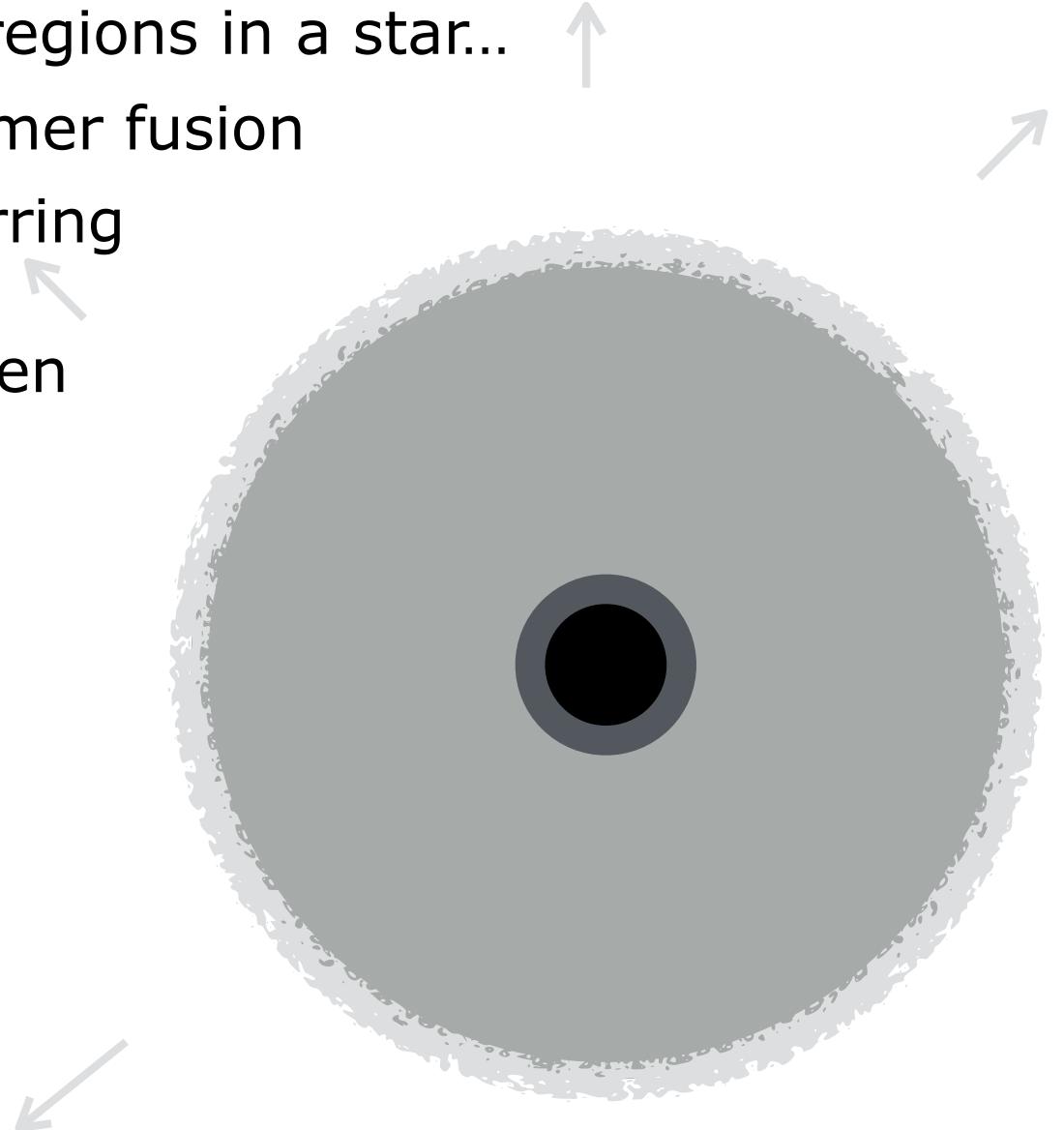


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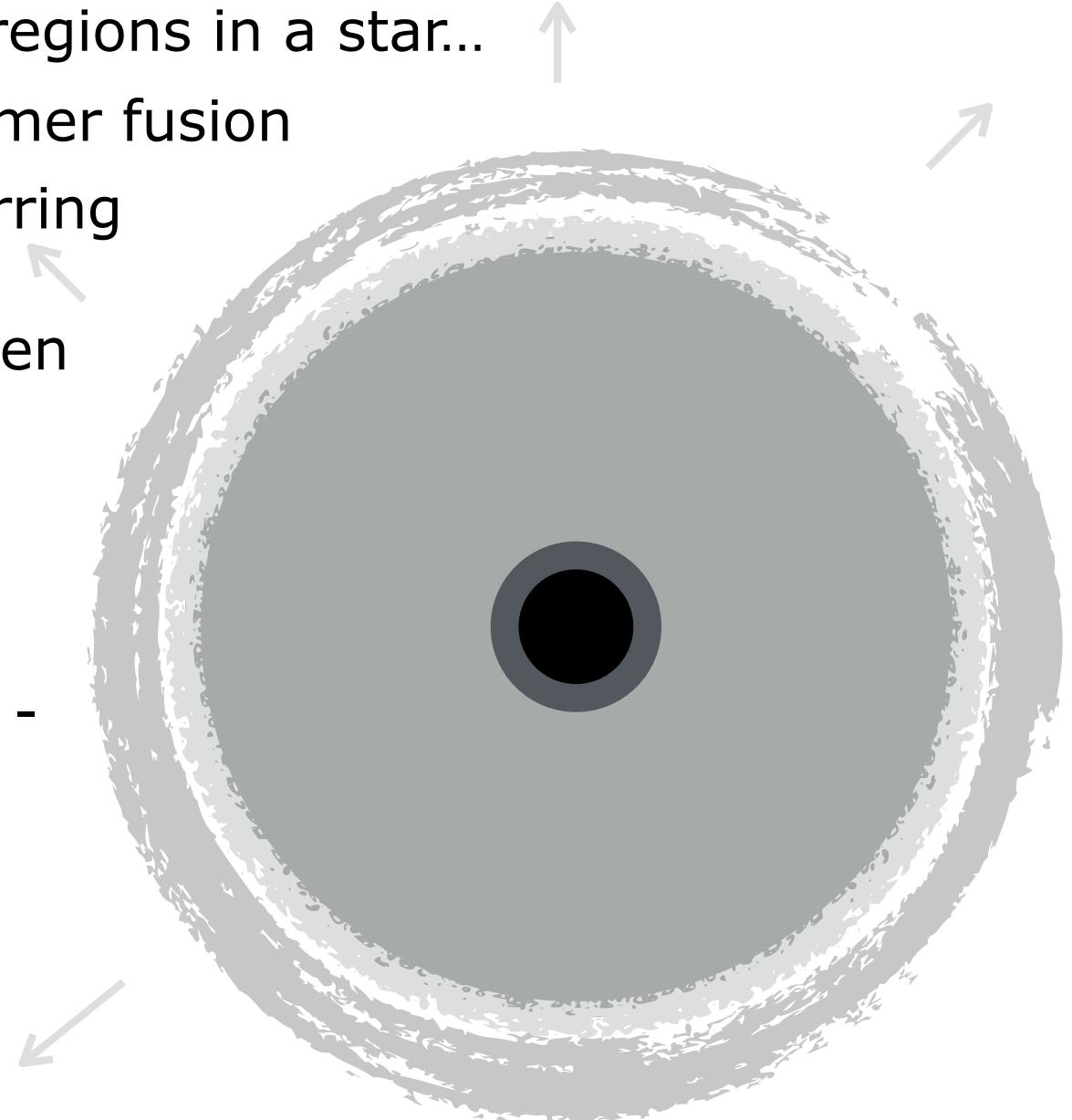


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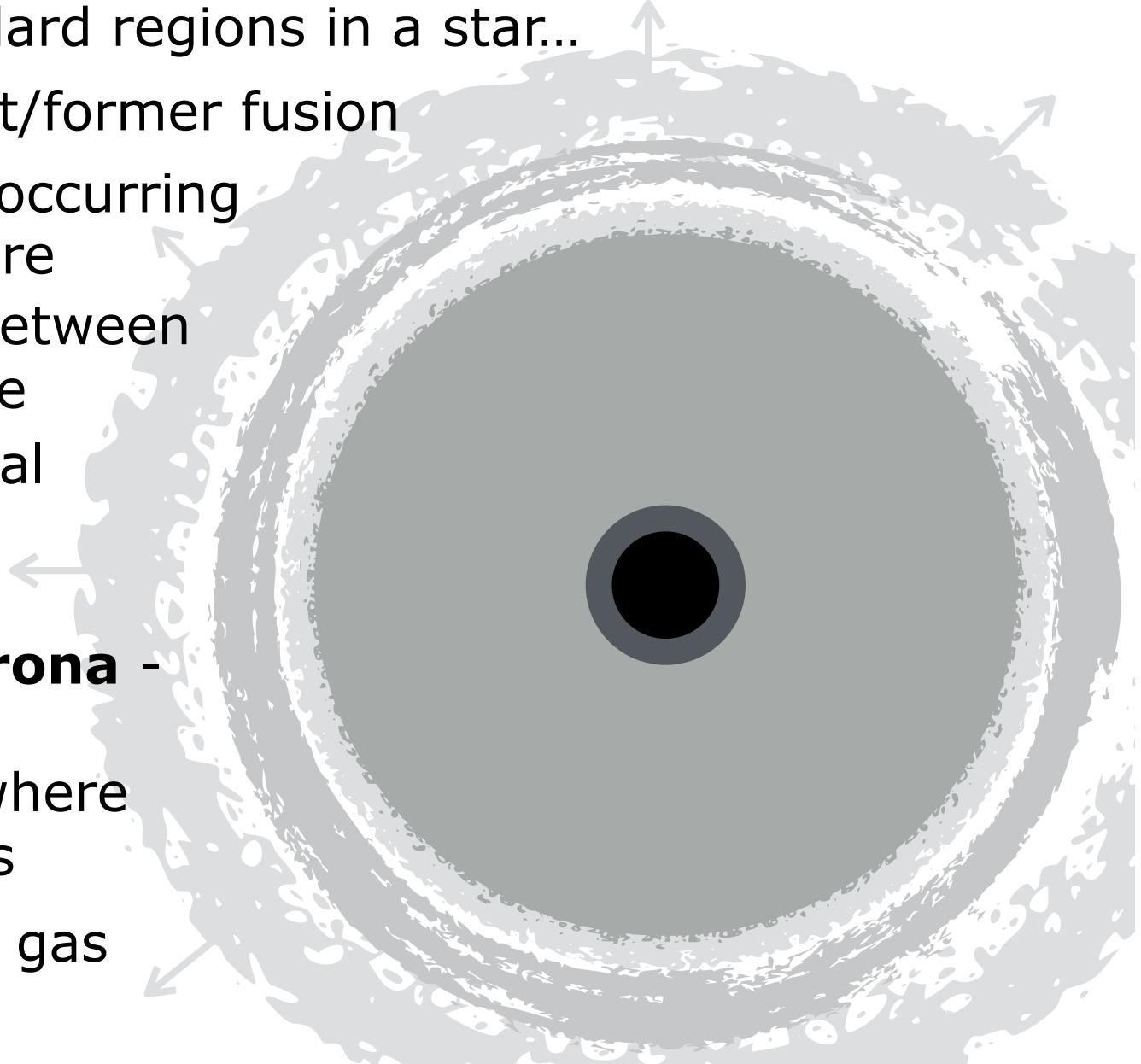


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- f) **wind** - region where gas escapes at $>> v_{\text{esc}}$



Stellar Evolution in a Nutshell

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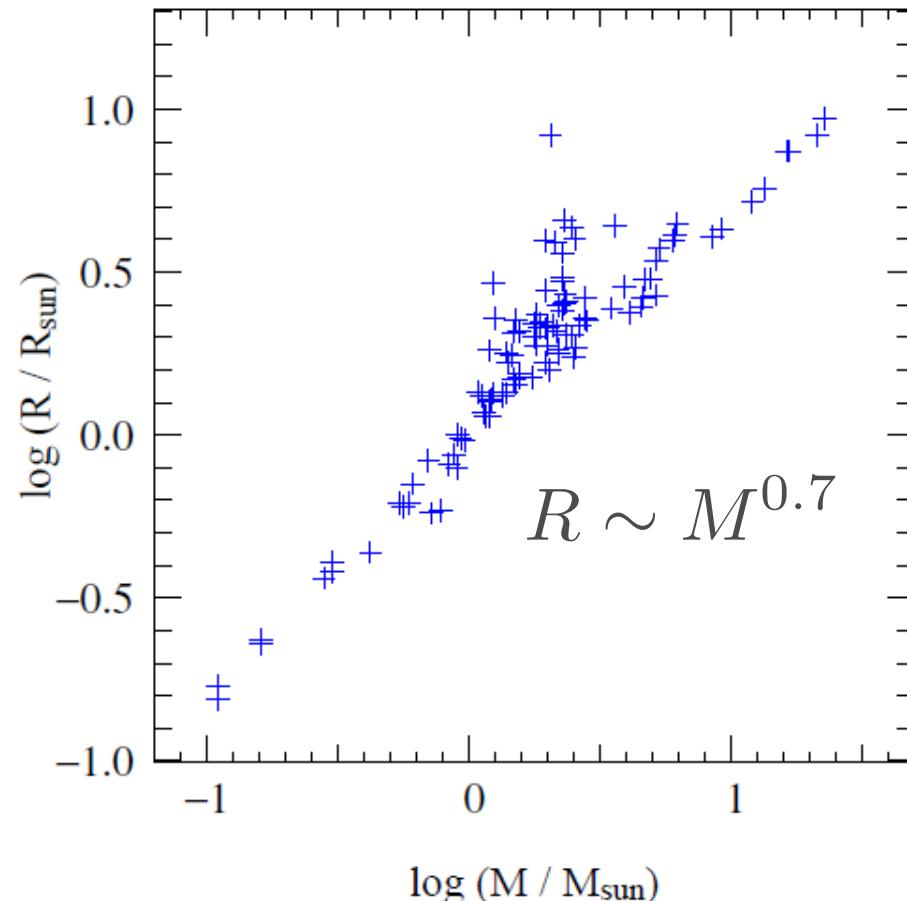
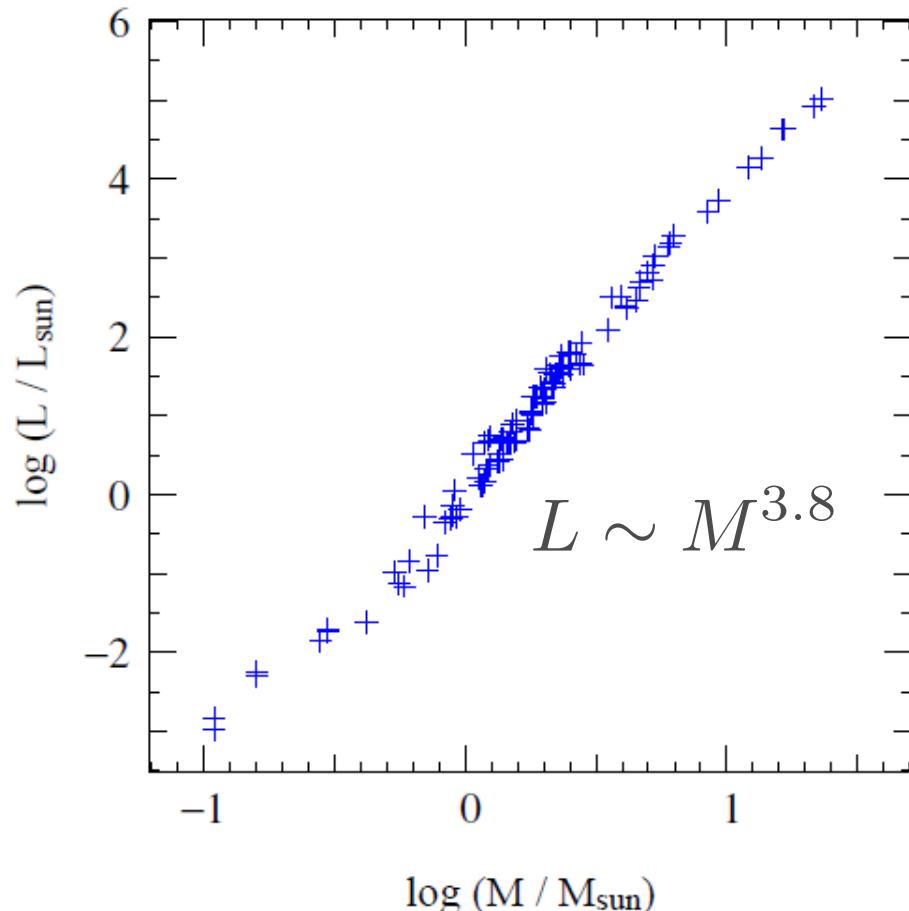
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- this cycle continues until fusion fuel is exhausted. Most stars stay in a perfect hydrostatic balance overall during this evolution.
- once the nuclear energy source is gone, gravity wins. In low-mass stars the core compresses and becomes a white dwarf. In high-mass stars the core collapses, producing a neutron star or black hole, and ejects the outer layers, producing a supernova.

Stellar properties relations

Mass-luminosity & Mass-radius relations

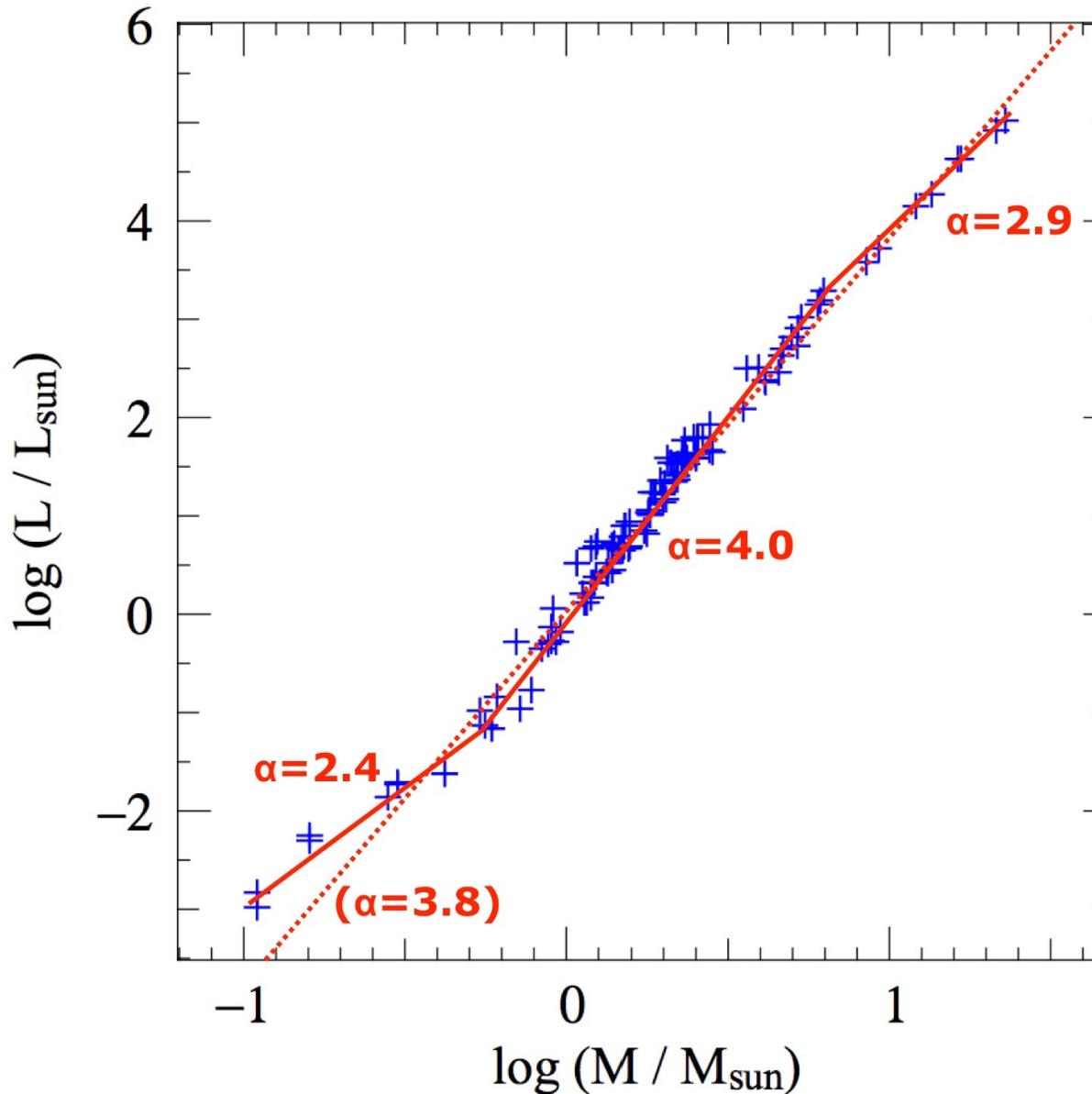
All 3 parameters are needed, so only possible for *main sequence stars** from double-lined eclipsing binaries

*stars that fuse H in their center



Stellar Evolution in a Nutshell

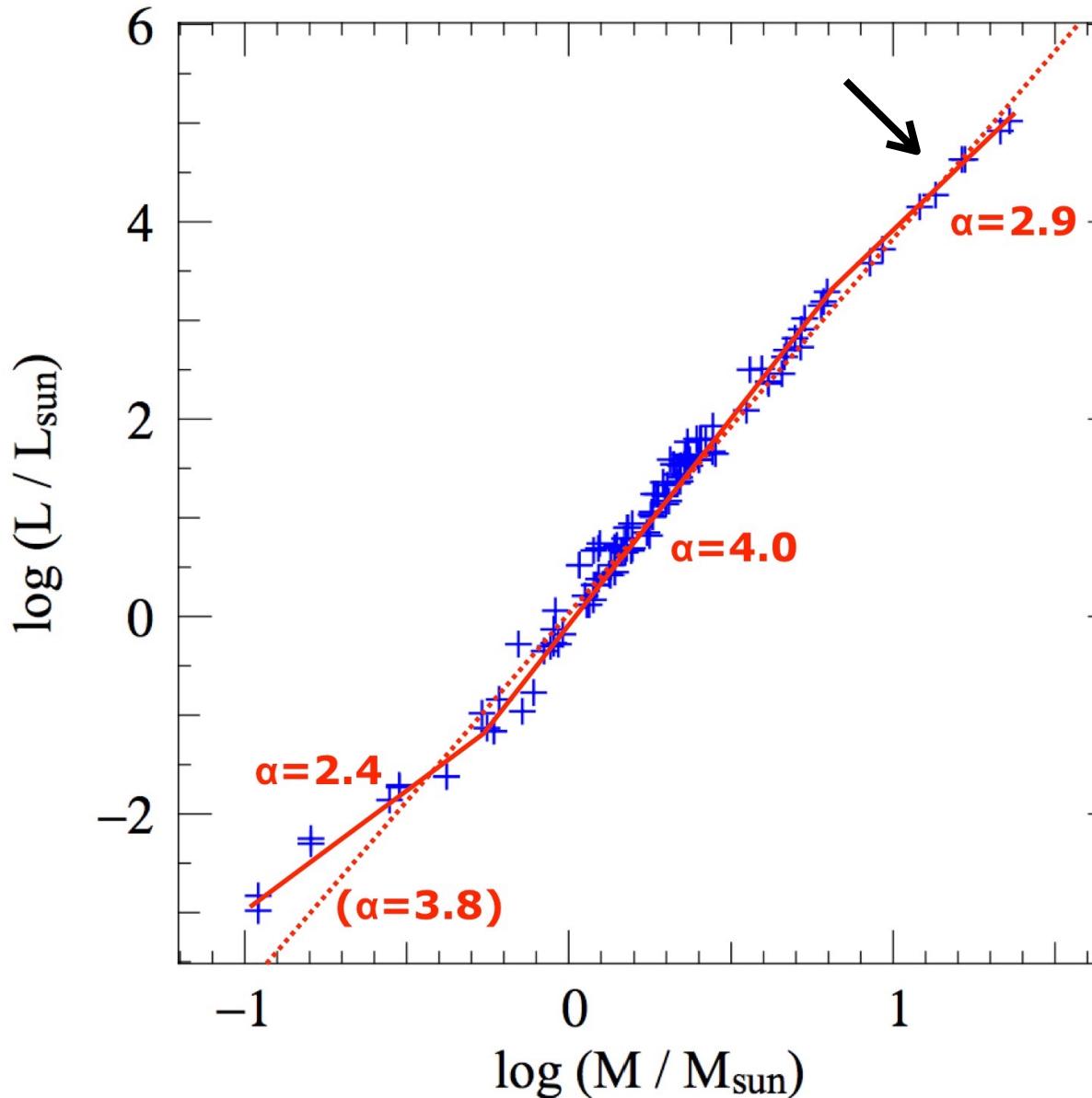
Mass-luminosity relation



Relation is only valid for stars during core H fusion (red giants/supergiants, degenerate stars, etc., do not follow the relation)

Stellar Evolution in a Nutshell

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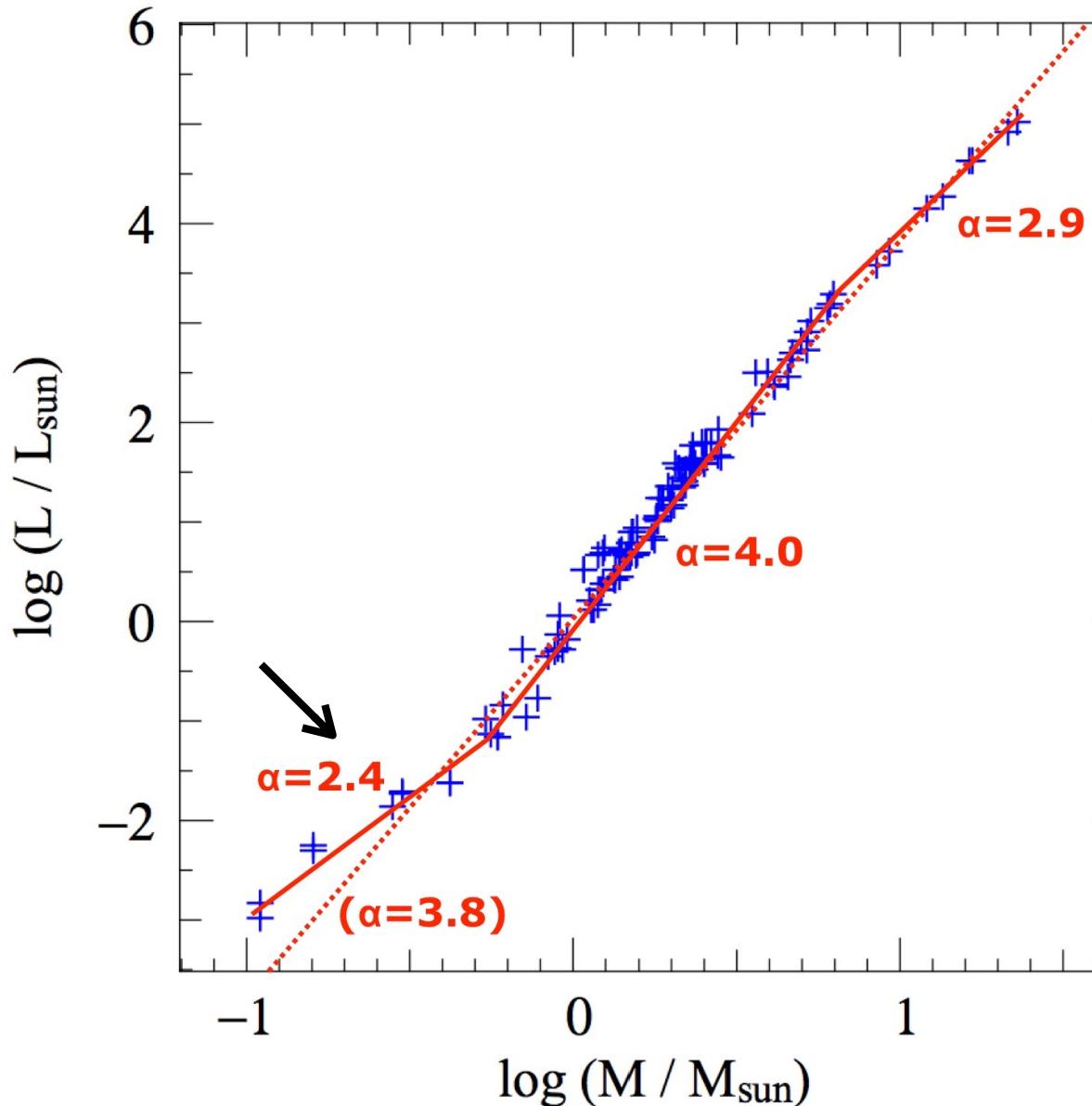


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Stellar Evolution in a Nutshell

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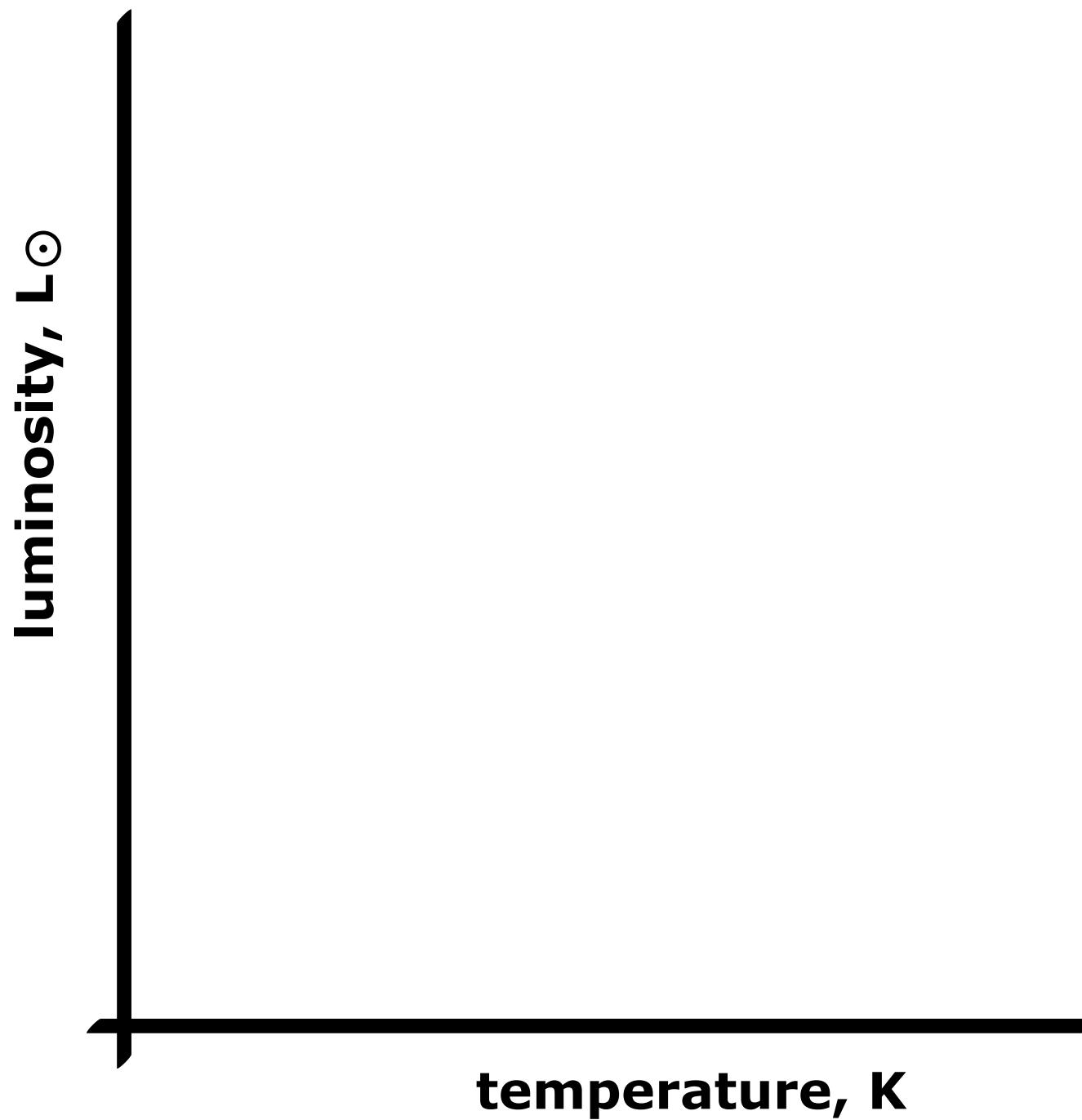


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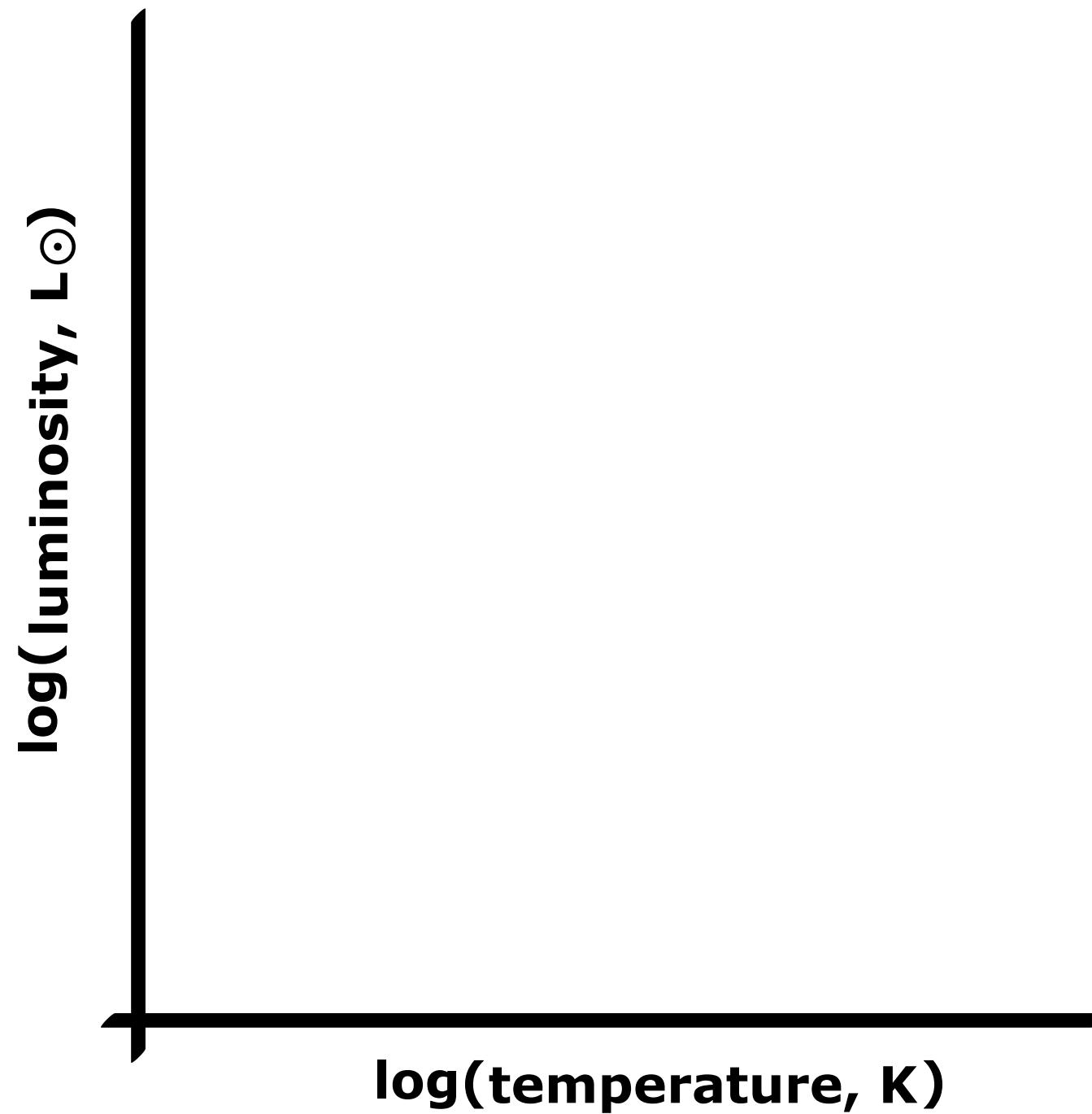
For $M > 6M_{\odot}$, slope flattens due to increasing role of radiation pressure

For $M < 6M_{\odot}$, slope flattens due to increasing role of convection throughout the star and changes in opacity

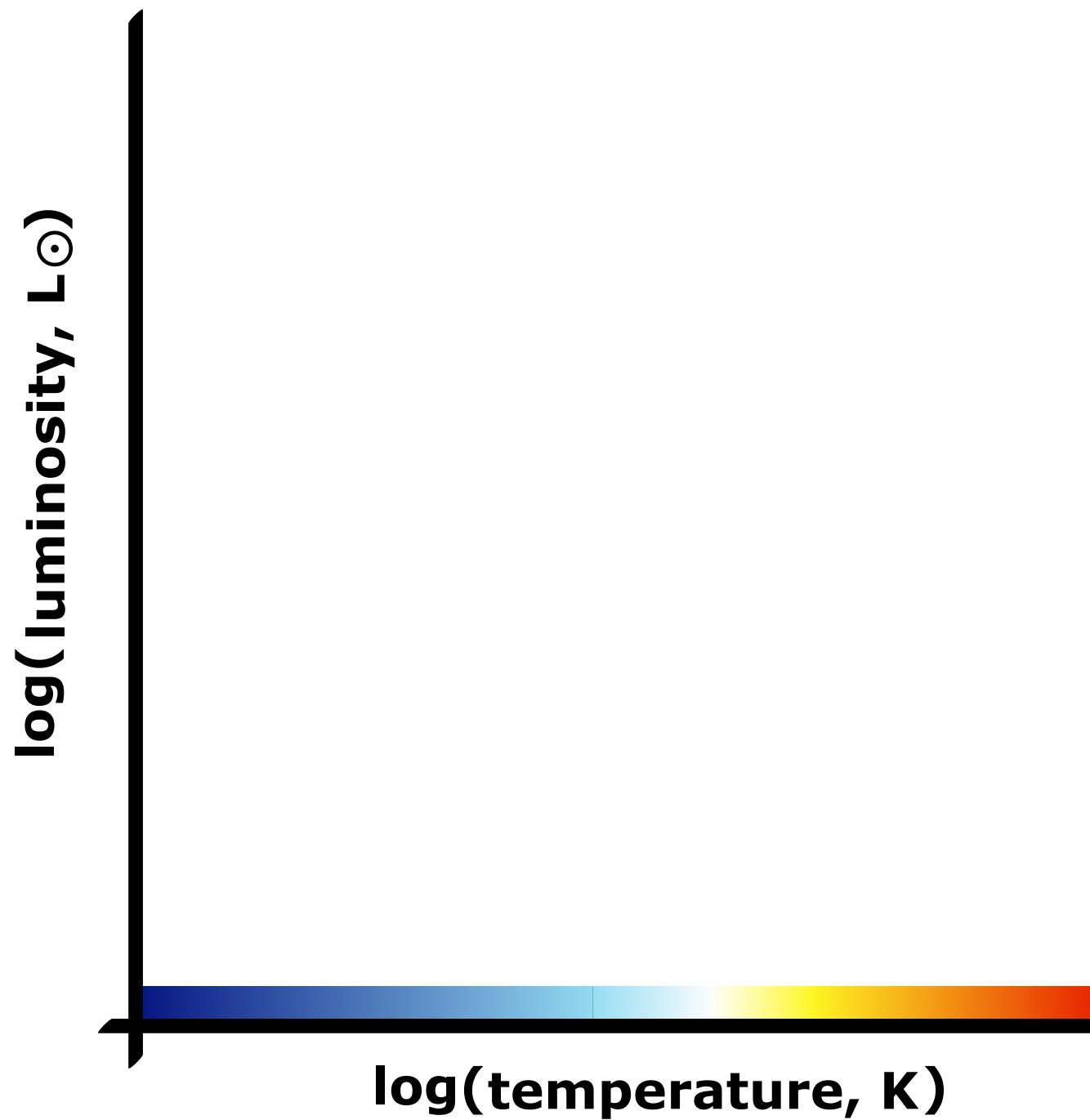
The Hertzsprung-Russell Diagram



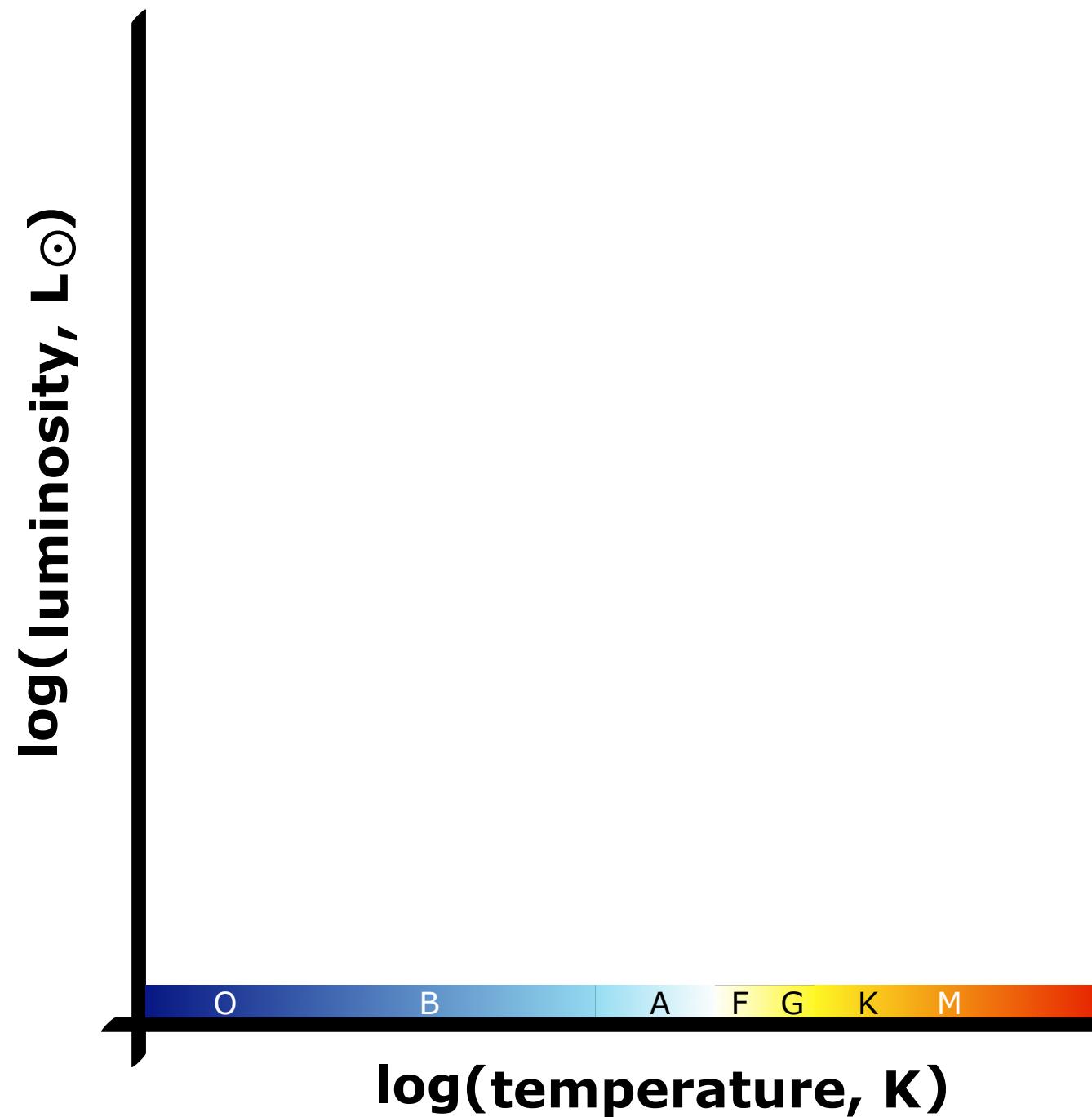
The Hertzsprung-Russell Diagram



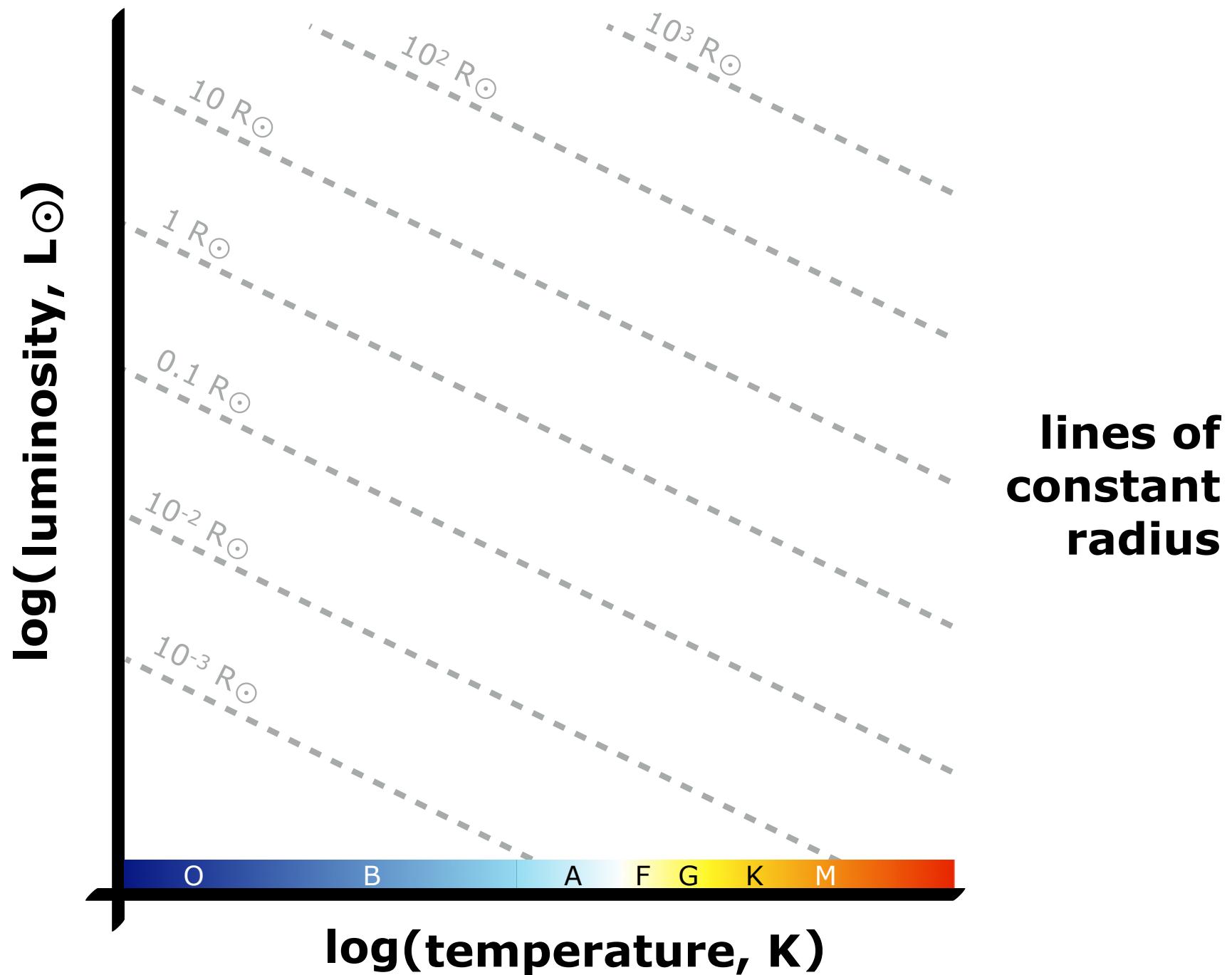
The Hertzsprung-Russell Diagram



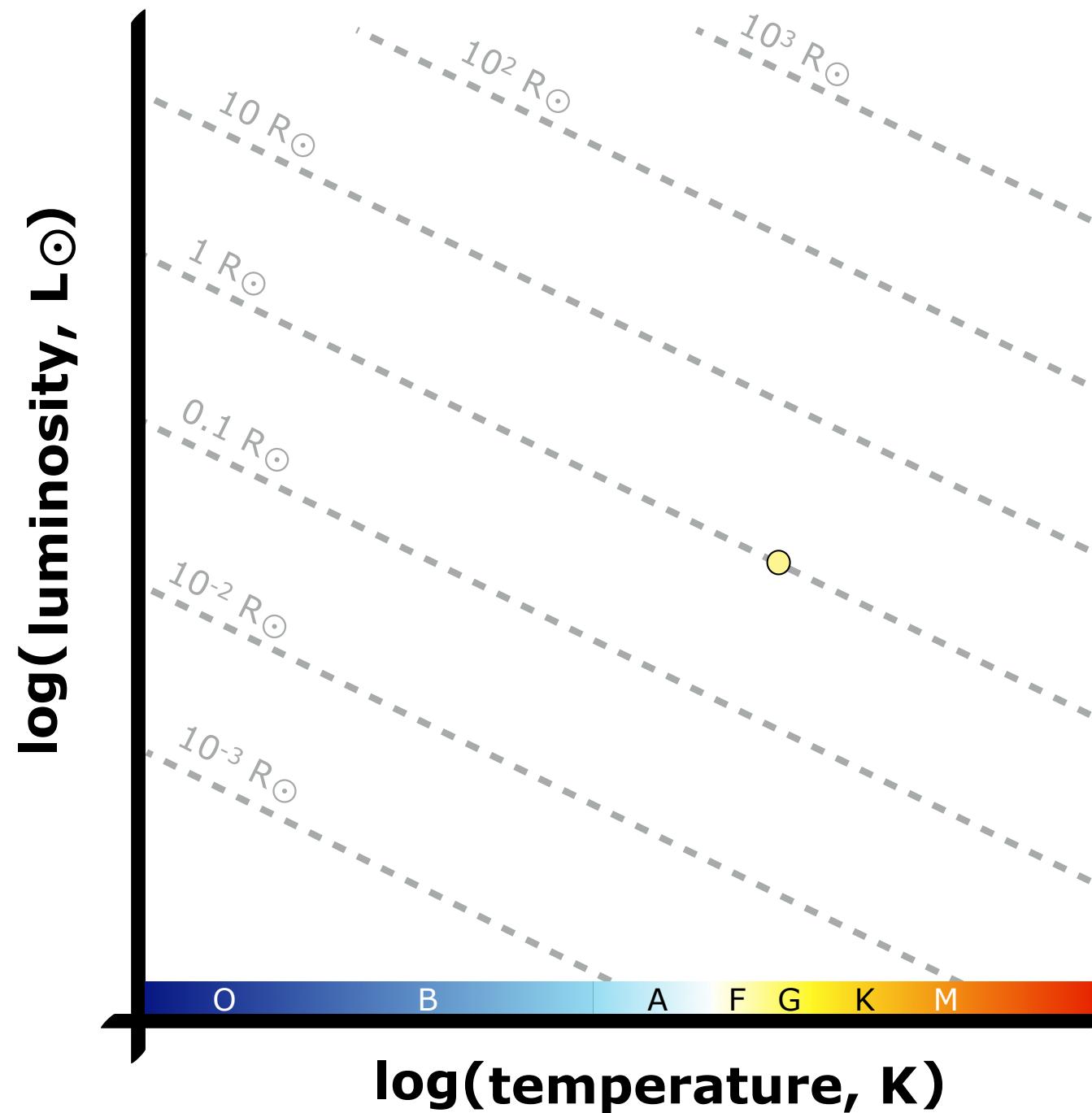
The Hertzsprung-Russell Diagram



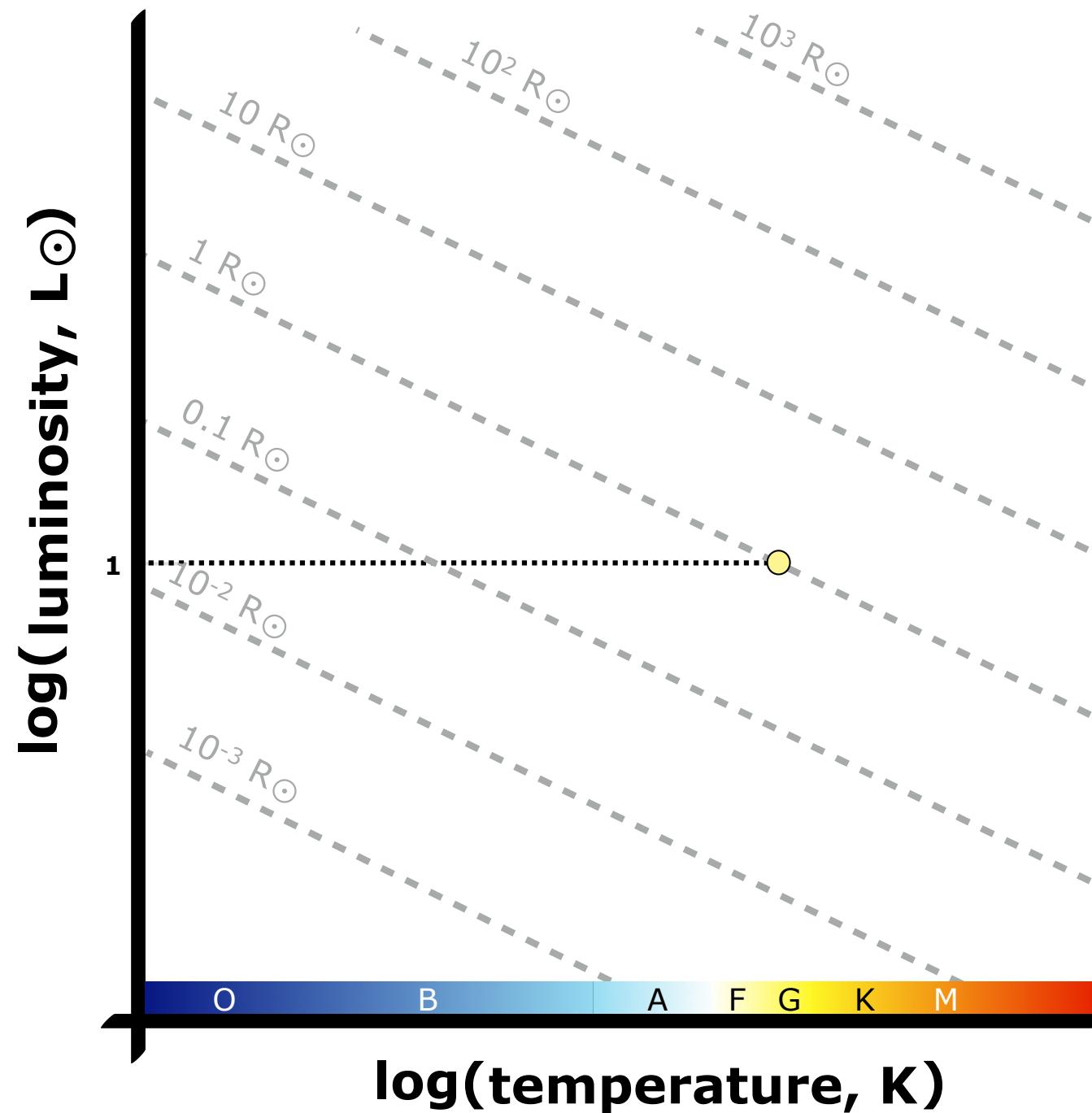
The Hertzsprung-Russell Diagram



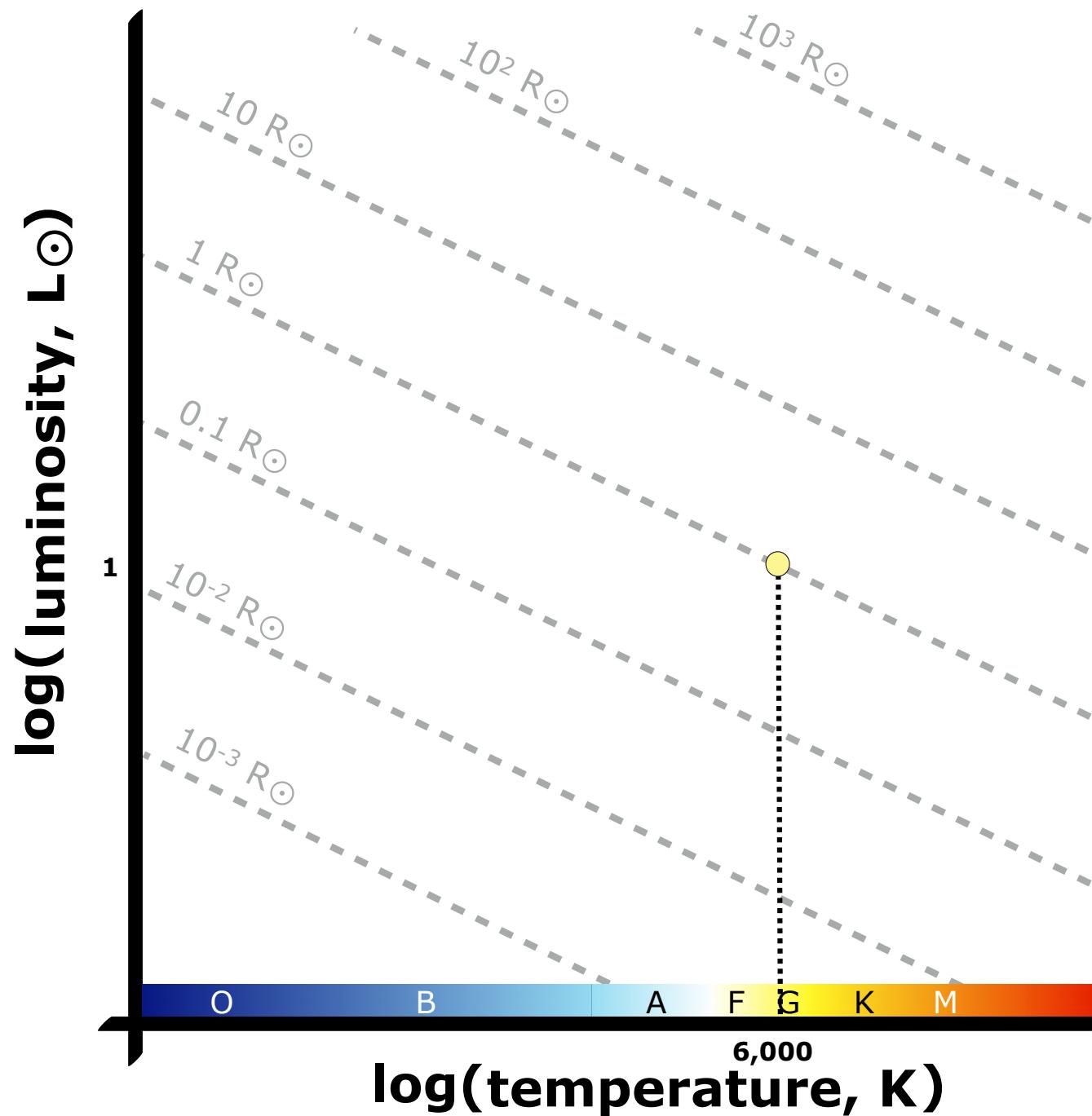
The Hertzsprung-Russell Diagram



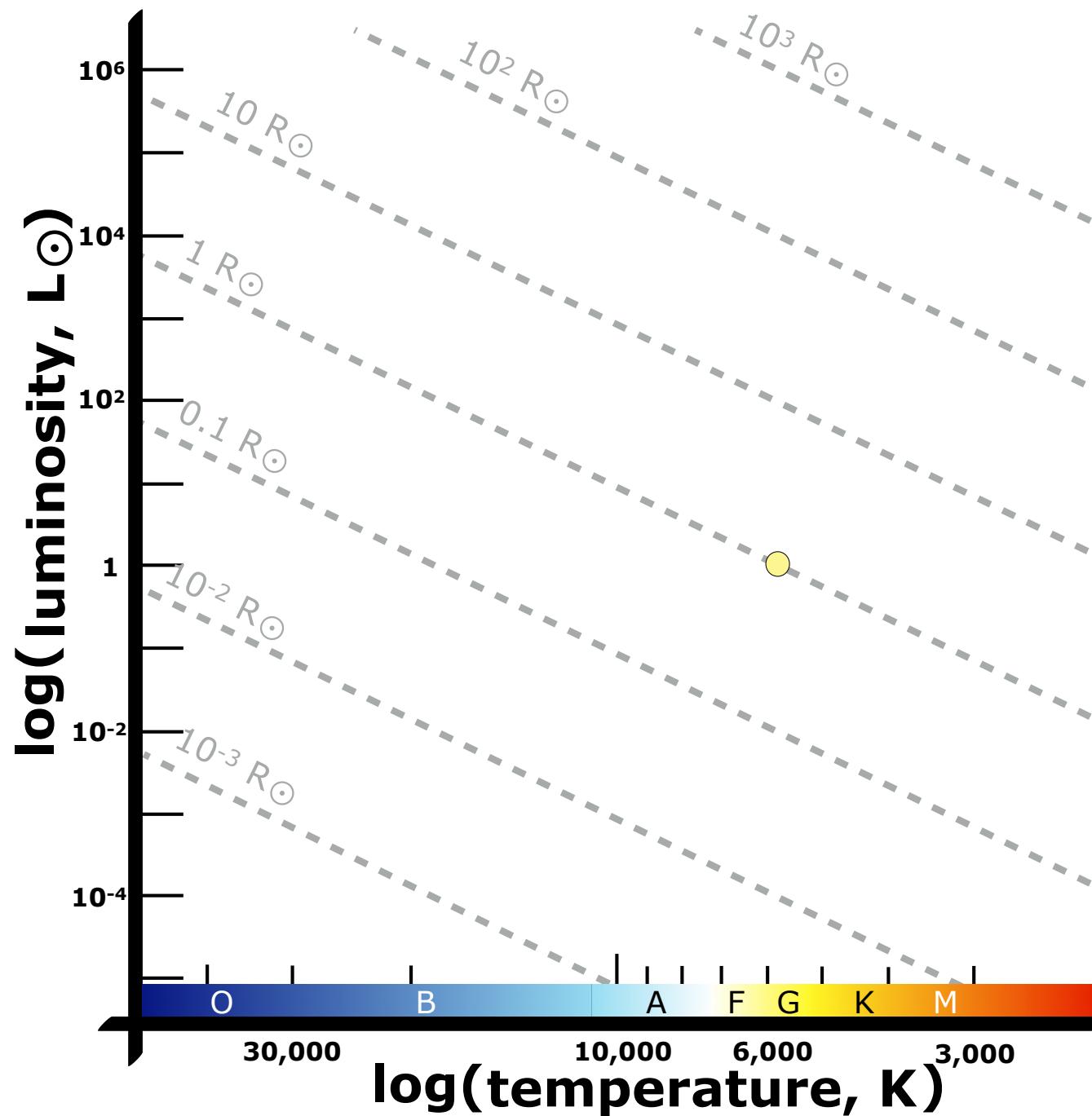
The Hertzsprung-Russell Diagram



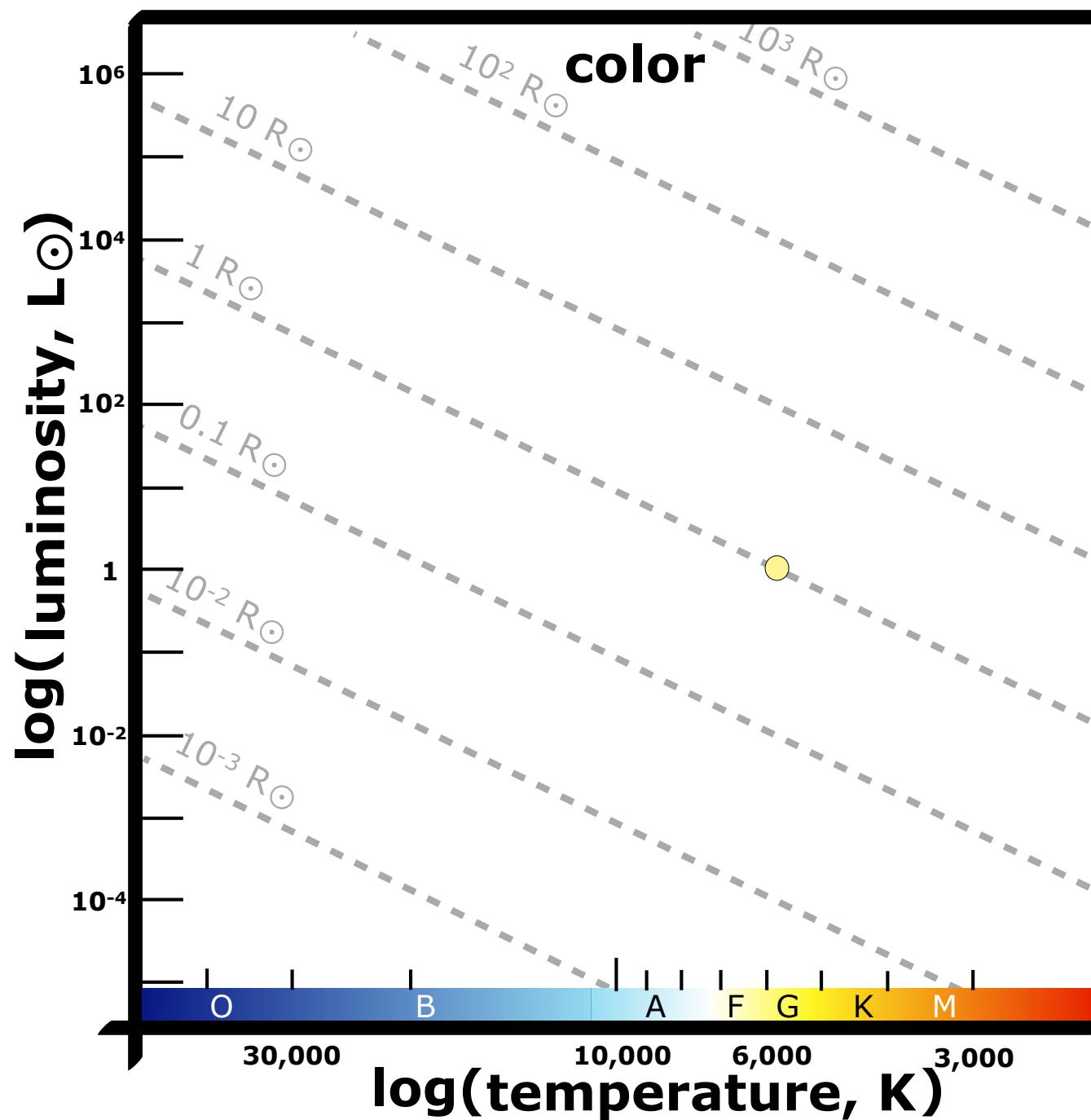
The Hertzsprung-Russell Diagram



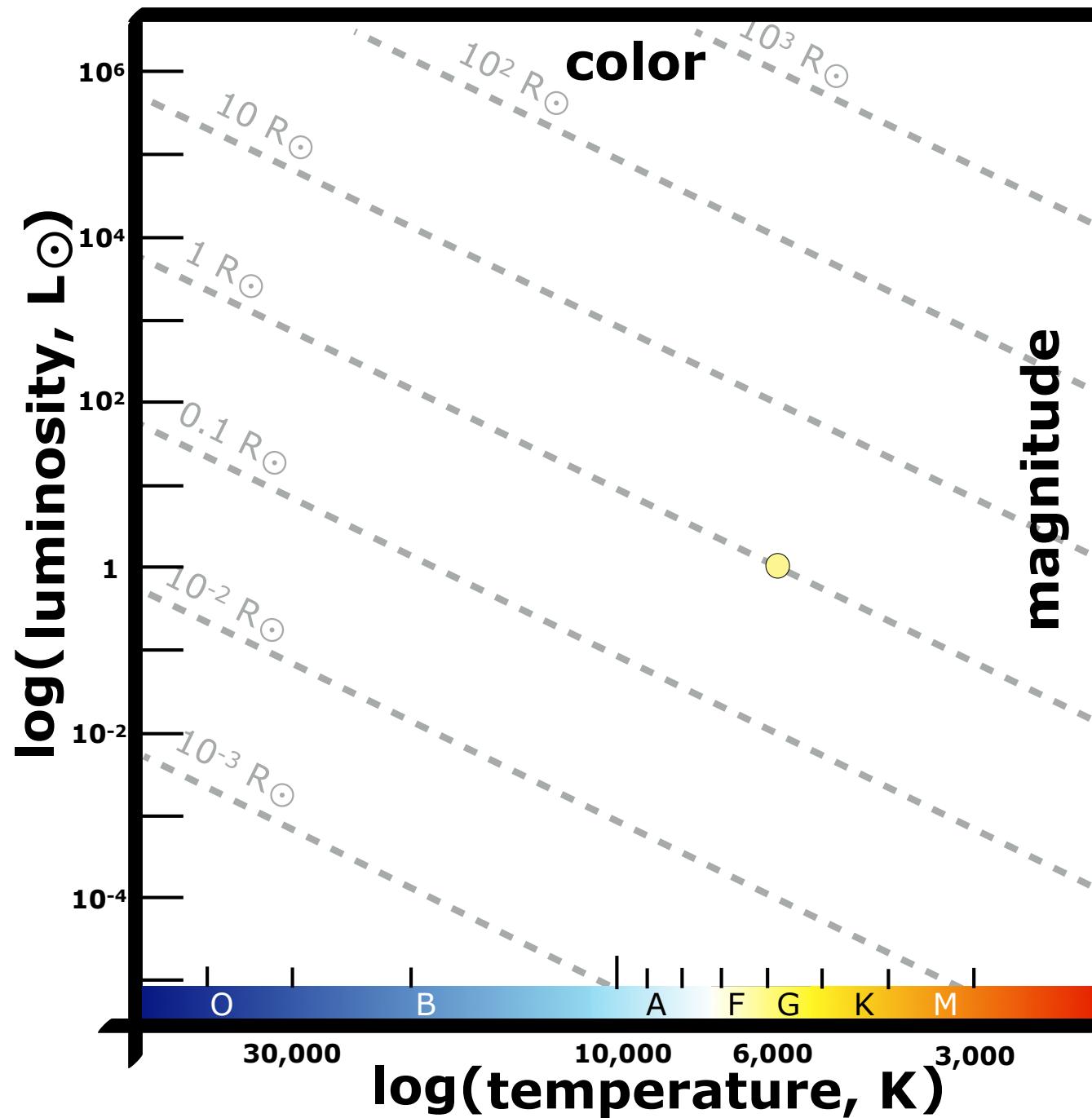
The Hertzsprung-Russell Diagram



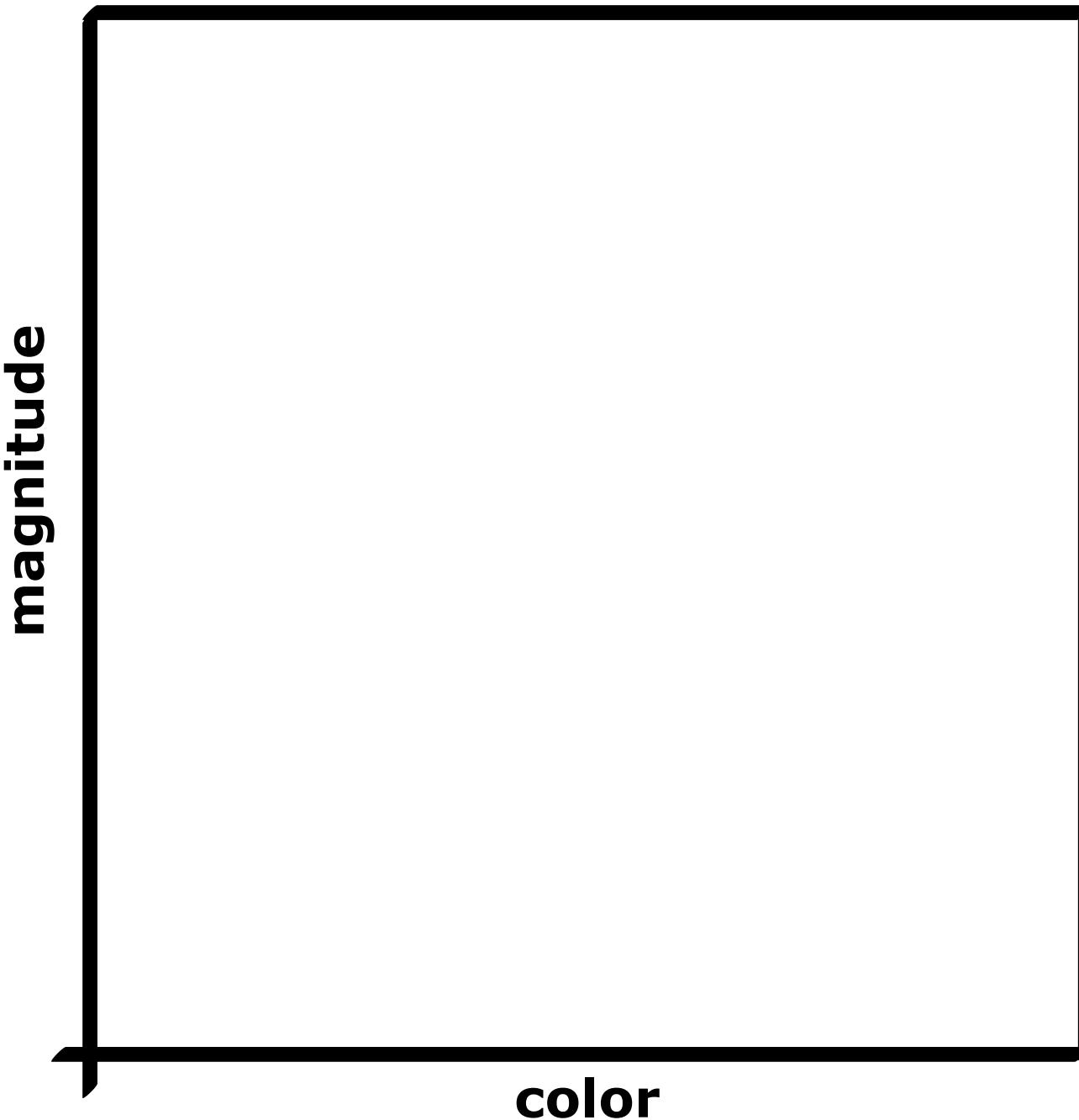
The Hertzsprung-Russell Diagram



The Hertzsprung-Russell Diagram



The Color-Magnitude Diagram



The Color-Magnitude Diagram

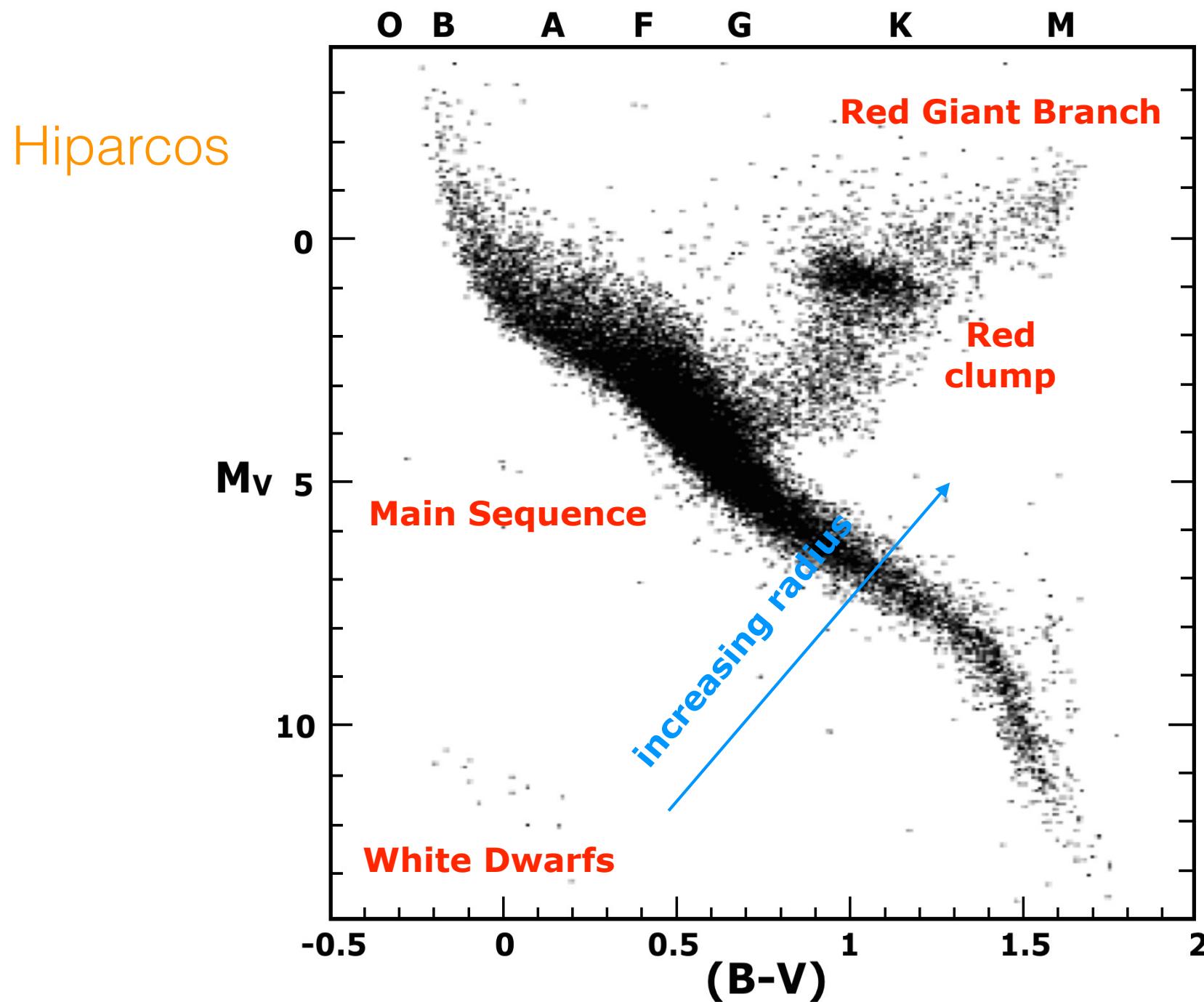
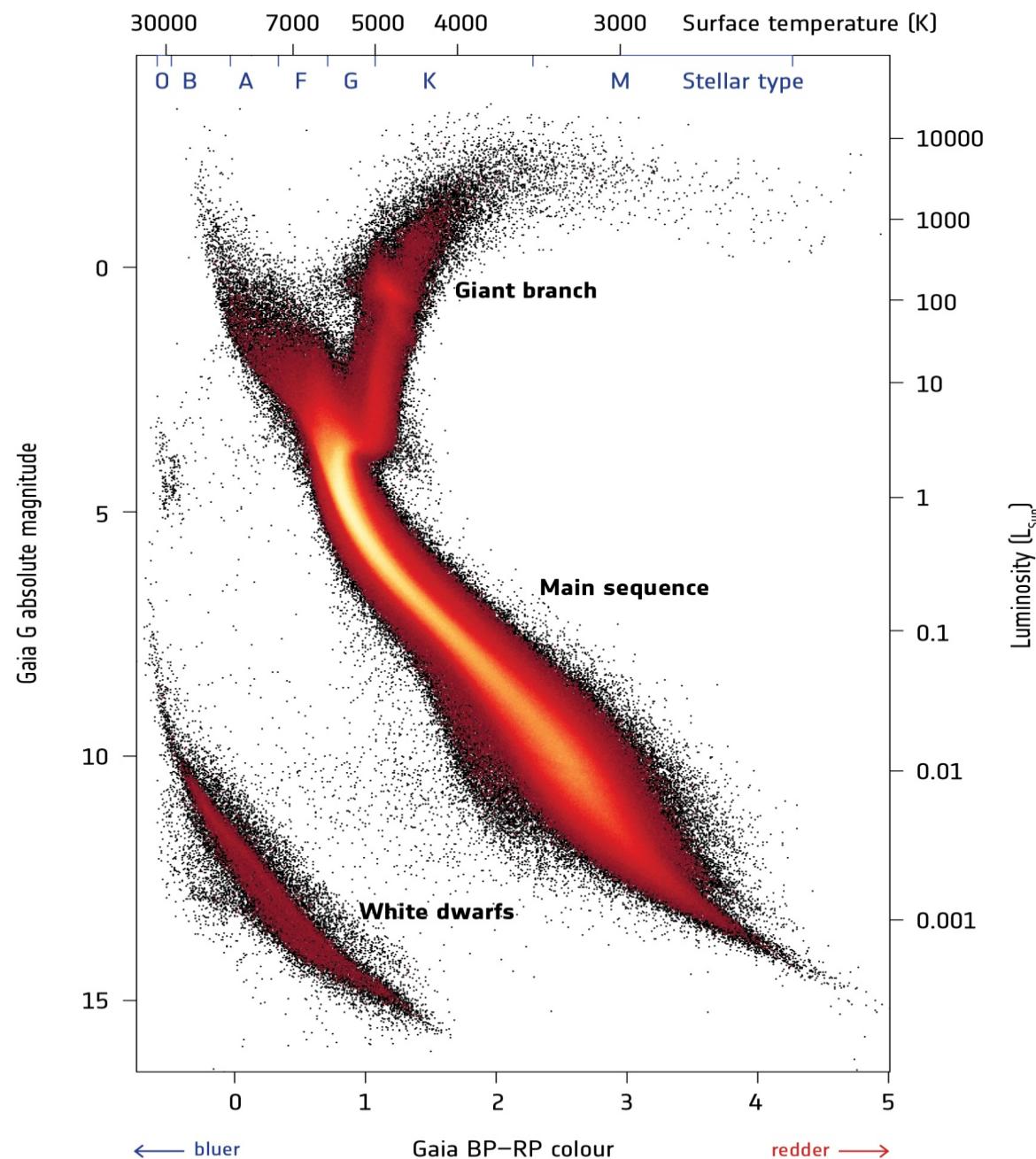


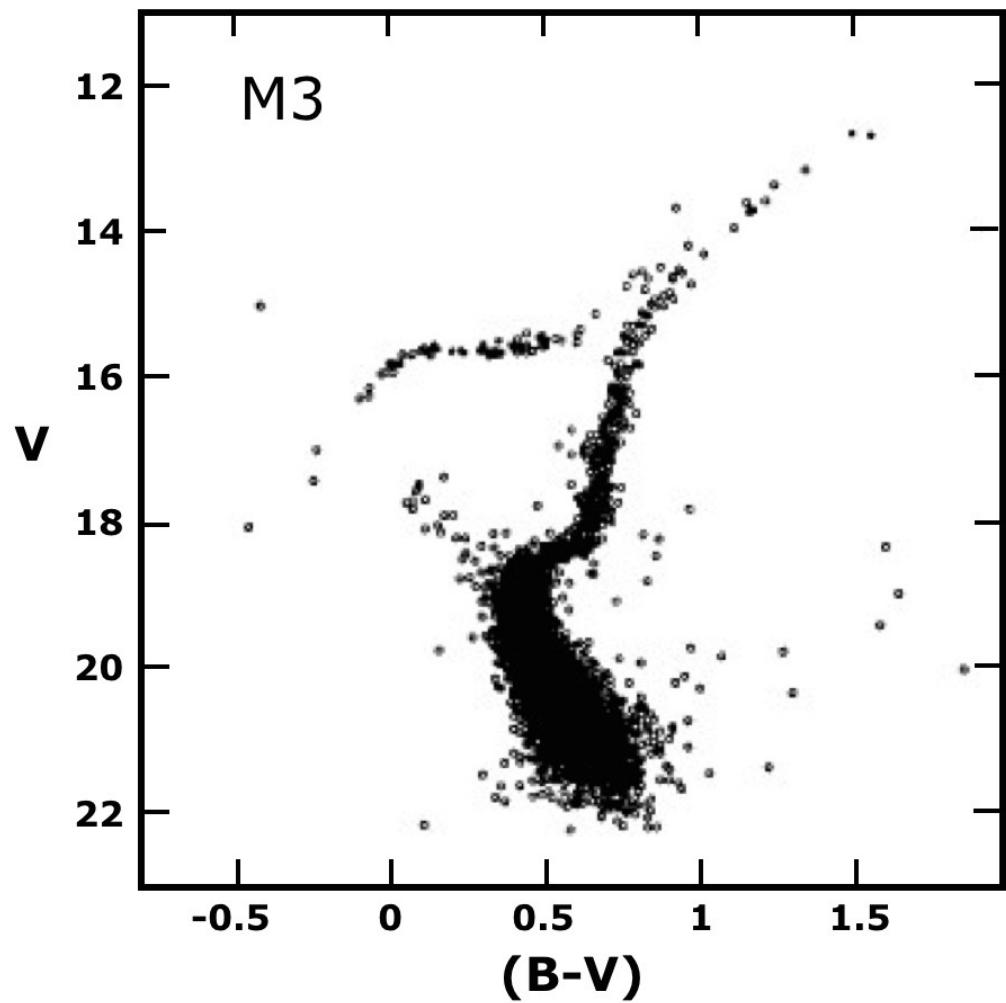
Figure courtesy of ESA.

The Color-Magnitude Diagram

Gaia



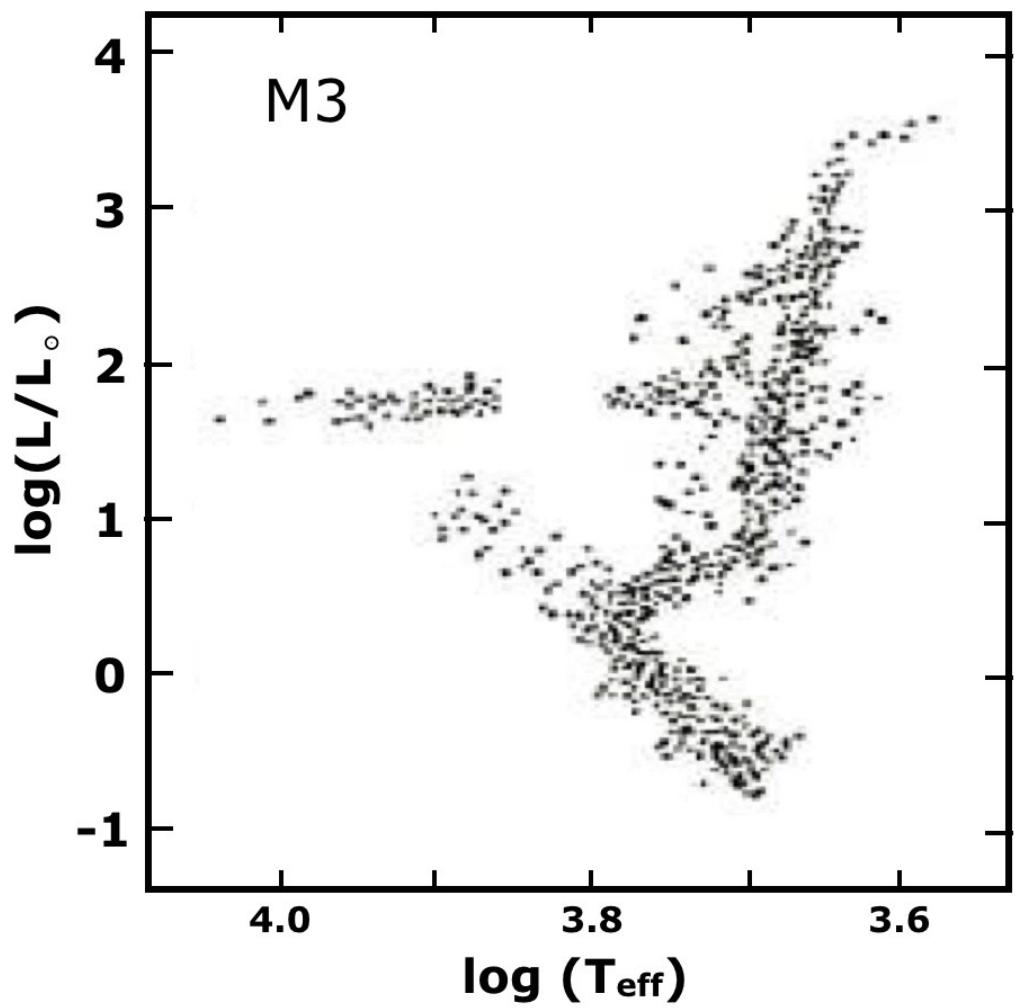
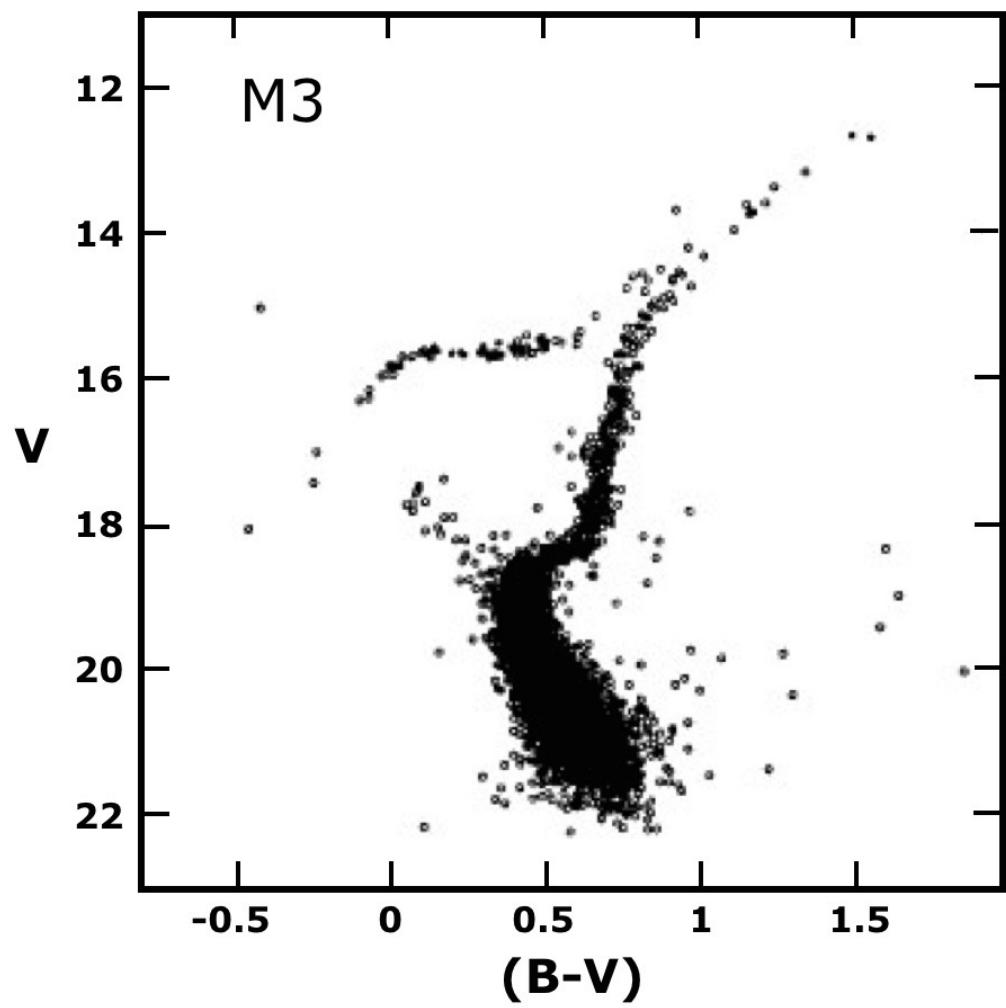
CMD and HRD



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Right: Reproduced from Johnson, H. L. and Sandage, A. R., 'Three-Color Photometry in the Globular Cluster M3', Astrophysical Journal, vol. 124, p.379, 1956.

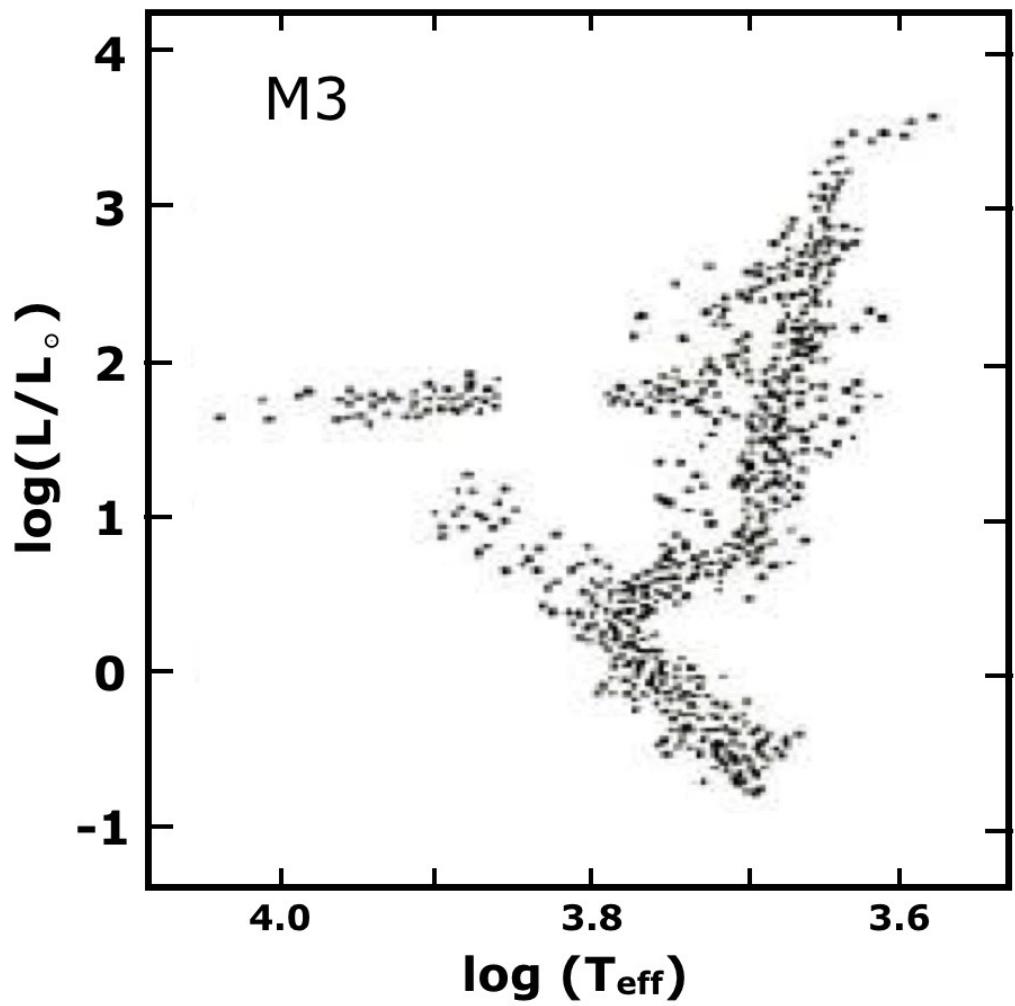
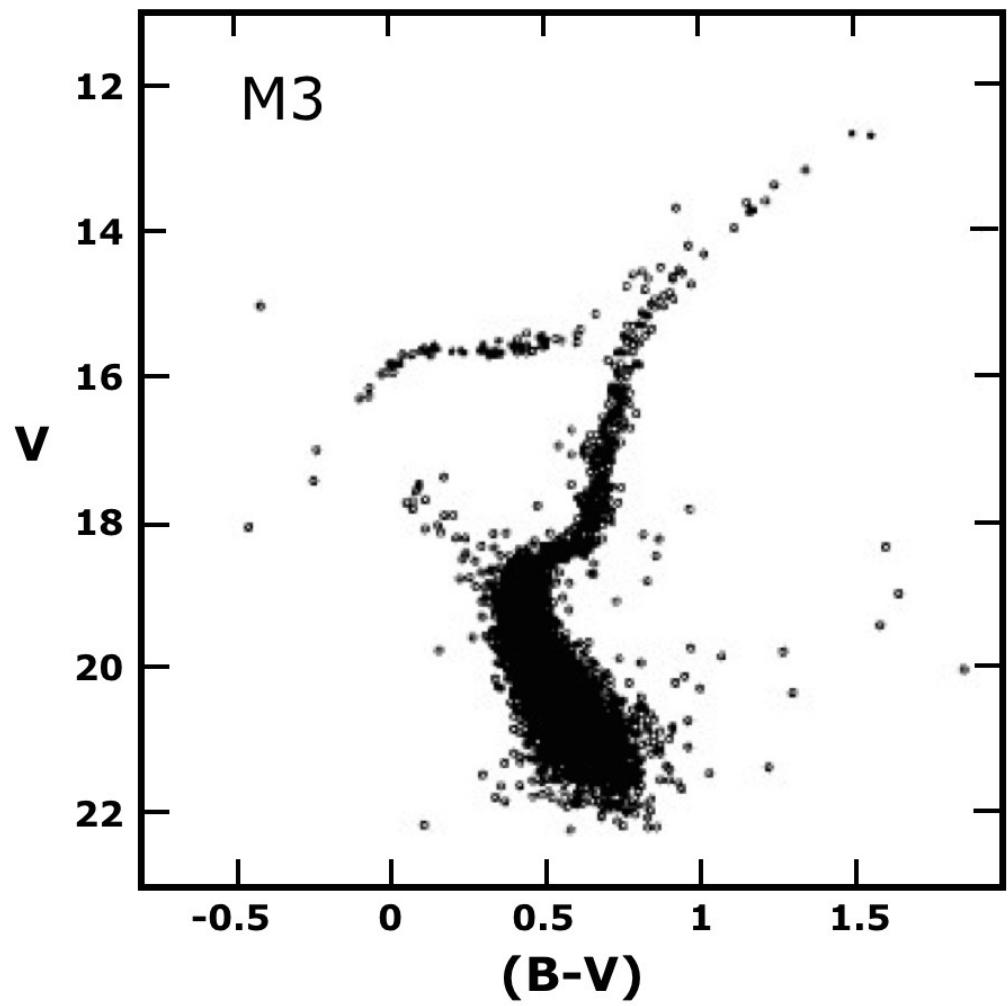
CMD and HRD



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CMD and HRD

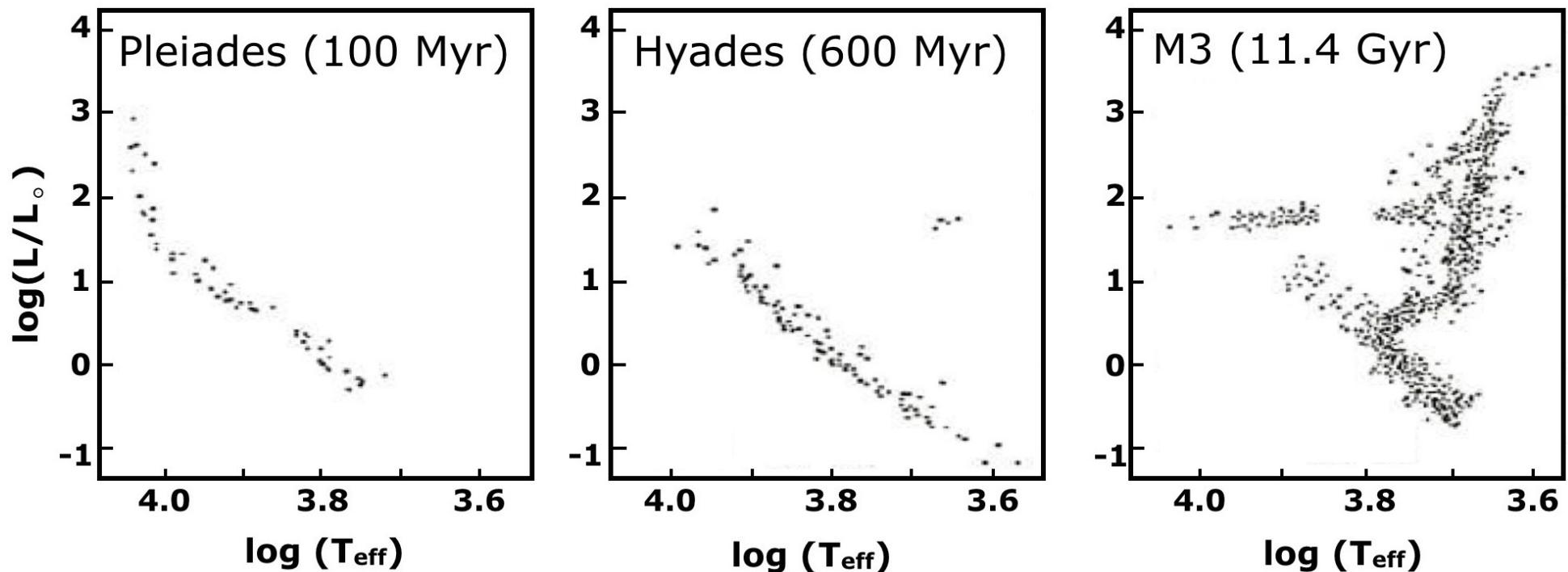


1. on the blue/hot side stars cover a small $(B-V)$, large T_{eff} range
2. on the red/cool side, stars cover a small T_{eff} , large $(B-V)$ range
3. horizontal branch curves down strongly in the CMD

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Right: Reproduced from Johnson, H. L. and Sandage, A. R., 'Three-Color Photometry in the Globular Cluster M3', Astrophysical Journal, vol. 124, p.379, 1956.

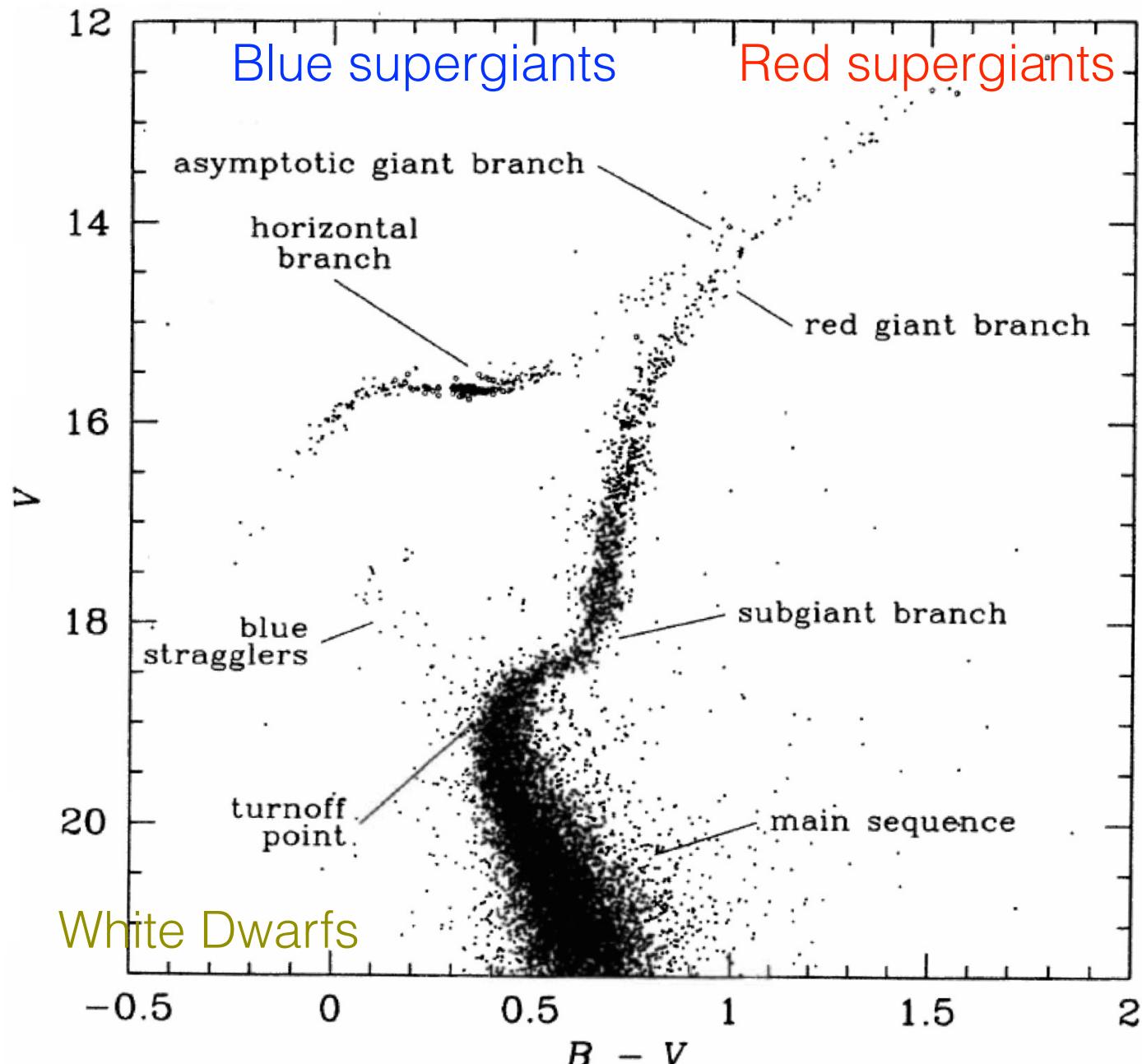
The Color-Magnitude Diagram



As the clusters increase in age:

- the MS gets shorter
- the ratio of red giants to MS stars increases

The Color-Magnitude Diagram



Reproduced from Alvio Renzini and Flavio Fusi Pecci, 'Tests of Evolutionary Sequences Using Color-Magnitude Diagrams of Globular Clusters', Annual Review of Astronomy and Astrophysics 1988 26:1, 199-244. Reproduced with permission.

Stellar populations



Population I: in the galactic disk, spiral arms and open clusters. $t < 10^9$ yr, $Z \sim 0.5 - 1 Z_\odot$

Population II: in the galactic disk, halo, globular clusters. $t \sim 10^{10}$ yr, $Z \sim 0.01 - 0.1 Z_\odot$

Population III: halo? $Z \sim 0?$

Basic assumptions: Th. Stell. Evo.

- Their structure and evolution depend only on intrinsic properties (mass and composition)
 - For binaries and dense clusters, evolution can be influenced by interaction
-
- Because molecular clouds are well-mixed
 - Assuming a quasi-solar composition: $X=0.706$ $Y=0.28$, $Z=0.014$
-
- Promoted by self-gravity.
 - Possible deviation from non-central forces: No rotation, no magnetic fields

Basic assumptions: Th. Stell. Evo.

1. Isolated in space

- Their structure and evolution depend only on intrinsic properties (mass and composition)
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Basic assumptions: Th. Stell. Evo.

1. Isolated in space

- Their structure and evolution depend only on intrinsic properties (mass and composition)
- For binaries and dense clusters, evolution can be influenced by interaction

2. Formed with a homogeneous composition

- Because molecular clouds are well-mixed
- Assuming a quasi-solar composition: $X=0.706$ $Y=0.28$, $Z=0.014$

- Promoted by self-gravity.
- Possible deviation from non-central forces: No rotation, no magnetic fields

Basic assumptions: Th. Stell. Evo.

1. Isolated in space

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- For binaries and dense clusters, evolution can be influenced by interaction

2. Formed with a homogeneous composition

- Because molecular clouds are well-mixed
- Assuming a quasi-solar composition: $X=0.706$ $Y=0.28$, $Z=0.014$

3. Have spherical symmetry

- Promoted by self-gravity.
- Possible deviation from non-central forces: No rotation, no magnetic fields

Exercises

1. *Stars are assumed to be isolated in space.* The star closest to the sun, Proxima Centauri, is 4.3 light-years away. How many solar radii is that? How it affects the Sun gravitationally? How this force compares to the attractive gravitational force the Sun suffers from Jupiter?
2. *Stars are assumed to form with a uniform composition.* What elements is the Sun made of?
3. The Sun rotates around its axis every 27 days. Calculate the ratio of the centrifugal acceleration a over the gravitational acceleration g for a mass element on the surface of the Sun.
4. The masses of stars are approximately in the range $0.08 M_{\odot} \leq M \leq 100 M_{\odot}$. Why is there an upper limit? Why is there a lower limit?
5. Can you think of methods to measure (1) the mass, (2) the radius, and (3) the luminosity of a star?
6. Think of a method to estimate the age of the clusters, discuss with your fellow students. Estimate the ages and compare with the results of your fellow students. Can you give an error range on your age estimates?

Exercises (2)

- 2.1 (a) Calculate the radius of an M5 I supergiant with $\log(L/L_\odot) = 5.50$ and $T_{\text{eff}} = 2700$ K.
(b) Assume a mass of approximately $20M_\odot$ and calculate the mean density of the star.
(c) Calculate the escape velocity.
(d) Compare these values with those for the Sun.
- 2.2 (a) Calculate the luminosities of the horizontal branch stars with $B-V \approx +0.30$ and $BC = 0.11$ in the cluster M3 (NGC 5272). The distance to M3 is 10.4 kpc and its interstellar extinction is negligible.
(b) What is the bolometric correction of the two stars at $B-V \approx 0.25$?
- 2.3 The star τ Sco has an apparent visual magnitude of $V = +2.8$ and a spectral type of approximately B0V. Parallax measurements indicate a distance of 470 ly.
(a) Calculate the absolute visual magnitude.
(b) Adopt the bolometric correction from Table 2.1 and calculate the luminosity L .
(c) Adopt the value of T_{eff} from Table 2.1 and calculate the radius.
(d) Estimate the mass.
(e) Calculate the acceleration of gravity at the stellar surface and the escape velocity.
(f) Calculate the mean density of the star.
(g) Compare these values with those of the Sun.
- 2.4 The Gaia satellite will measure parallaxes with an accuracy of 2×10^{-5} arcsec for stars with $V < 15$
(a) What is the distance d of an M5 I supergiant with $\log(L/L_\odot) = 5.50$ and $BC = 3.70$ that has $V = 12$, if the effect of interstellar extinction is ignored?
(b) What is the relative distance accuracy $\sigma(d)/d$ of such a star?
(c) Assuming that T_{eff} is known, what is the relative accuracy in radius $\sigma(R)/R$?