DRAFT VERSION MAY 7, 2023
Typeset using LaTeX default style in AASTeX63

Erratum: "Kepler and the Behemoth: Three Mini-Neptunes in a 40 Million Year Old Association" (2022, AJ, 164, 215)

L. G. BOUMA,[1,*] R. KERR,[2] J. L. CURTIS,[3] H. ISAACSON,[4] L. A. HILLENBRAND,[1] A. W. HOWARD,[1] A. L. KRAUS,[2] A. BIERYLA,[5] D. W. LATHAM,[5] E. A. PETIGURA,[6] AND D. HUBER[7]

[1]*Cahill Center for Astrophysics, California Institute of Technology, Pasadena, CA 91125, USA*
[2]*Department of Astronomy, The University of Texas at Austin, Austin, TX 78712, USA*
[3]*Department of Astronomy, Columbia University, 550 West 120th Street, New York, NY 10027, USA*
[4]*Astronomy Department, University of California, Berkeley, CA 94720, USA*
[5]*Center for Astrophysics | Harvard & Smithsonian, 60 Garden St, Cambridge, MA 02138, USA*
[6]*Department of Physics & Astronomy, University of California Los Angeles, Los Angeles, CA 90095, USA*
[7]*Institute for Astronomy, University of Hawai'i, 2680 Woodlawn Drive, Honolulu, HI 96822, USA*

Submitted to Astronomical Journal

In the original manuscript, the top-right panel of Figure 2 erroneously omitted KOI-7913 B due to an error in the plotting script. Figure 1 in this erratum corrects the omission.

Corresponding author: L. G. Bouma
luke@astro.caltech.edu

* 51 Pegasi b Fellow

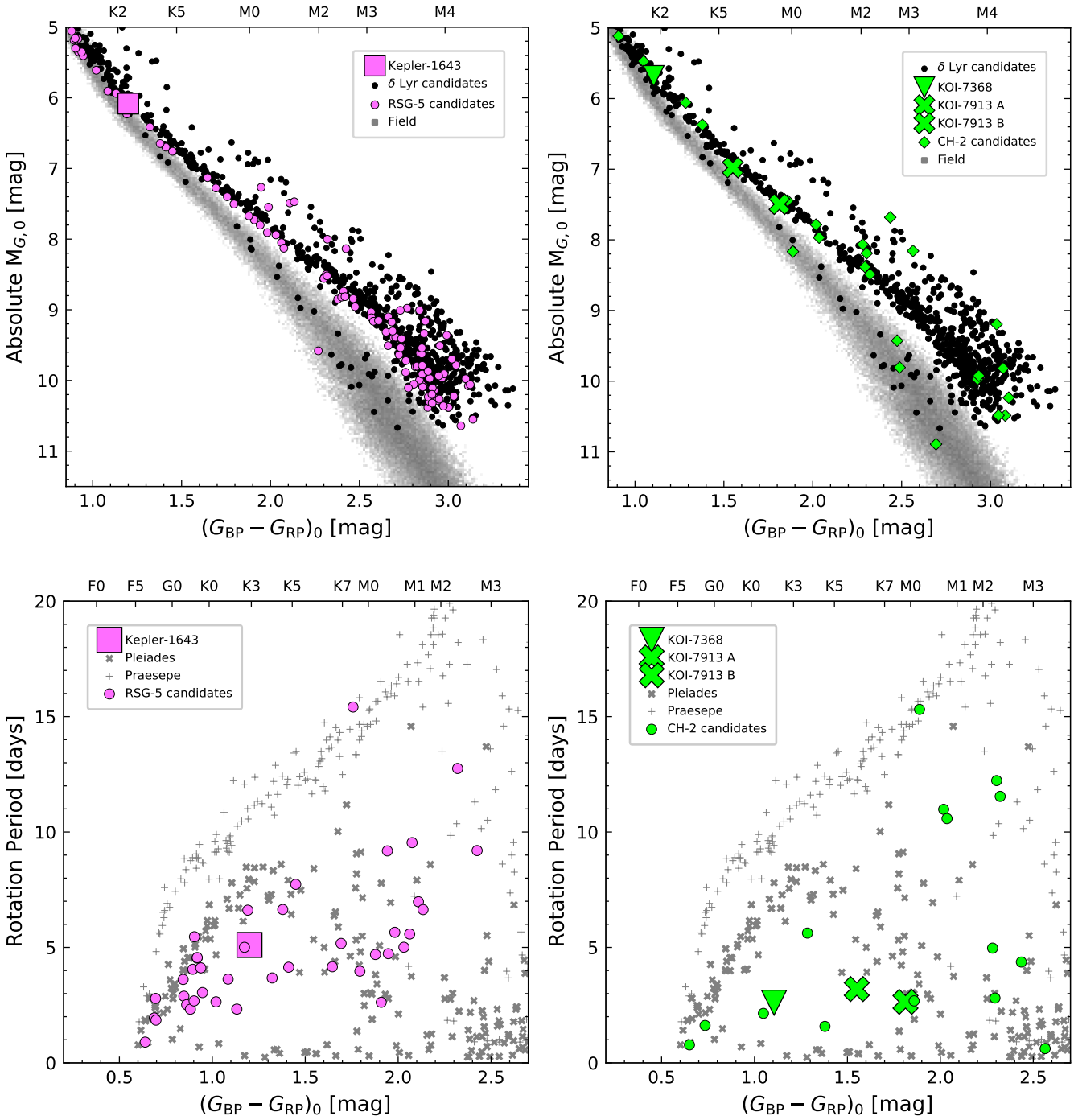

Figure 1. Age-diagnostic diagrams from the stellar groups near Kepler-1643, KOI-7368, and KOI-7913. *Top row*: Color–absolute magnitude diagram of candidate Cep-Her members, plotted over candidate members of the δ Lyr cluster (≈ 38 Myr; Bouma et al. 2022) and the Gaia EDR3 Catalog of Nearby Stars (gray background). The left and right columns shows stars in RSG-5 and CH-2, respectively. The range of colors is truncated to emphasize the pre-main-sequence; approximate spectral types are shown on the upper axes. Stars that fall far below the cluster sequences are field interlopers. *Bottom row*: TESS and ZTF-derived stellar rotation periods, with the Pleiades (≈ 112 Myr) and Praesepe (≈ 650 Myr) shown for reference (Rebull et al. 2016; Douglas et al. 2017). The detection efficiency for reliable rotation periods falls off beyond $(G_{BP} - G_{RP})_0 \gtrsim 2.6$.

REFERENCES

15 Bouma, L. G., Curtis, J. L., Masuda, K., et al. 2022, AJ, 163, 121

16 Douglas, S. T., Agüeros, M. A., Covey, K. R., & Kraus, A. 2017,
17 ApJ, 842, 83

18 Rebull, L. M., Stauffer, J. R., Bouvier, J., et al. 2016, AJ, 152, 113