

# For positioning electrodes in the ATLAS experiment's New Small Wheels using detector characterization techniques

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## **Abstract**

ATLAS is a multi-purpose particle detector system designed to capture the outcome of proton-proton collisions at the Large Hadron Collider (LHC) at CERN. The innermost end-caps of the ATLAS muon spectrometer consist of two wheels of muon detectors that must be replaced to improve the angular resolution of tracks for precision muon momentum reconstruction in the next phase of LHC operation. The New Small Wheels (NSWs) will be covered with two detector types that must trigger on and track outgoing particles - one type is small-strip thin gap chambers (sTGCs). Canada is responsible for one quarter of the required sTGCs. At McGill University, modules with four layers of sTGCs (called quadruplets) are characterized using a cosmic ray hodoscope before being sent to CERN for further testing and integration into the wheels. Quadruplets must be able to reconstruct particle tracks with 1 mrad angular resolution. Misalignments between sTGC layers must be corrected for to achieve this goal. The charge profile left by an x-ray gun and coordinate measuring machine (CMM) measurements of quadruplet layers are being used to define these parameters. Work on using cosmic ray data to validate misalignment parameters derived using the above-mentioned methods will be presented.

## Résumé

C'est le résumé.

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## **Acknowledgements**

Something along the lines of . . . I would like to thank all the little people who made this thesis possible.

### **Contribution of authors**

Something along the lines of . . . I hereby declare that I am the sole author of this thesis. This is a true copy of the thesis, including any required final revisions, as accepted by my examiners.

# Chapter 1

## Introduction

The High-Luminosity Large Hadron Collider (HL-LHC) project [1] was approved to combat the plateau in statistical gain of recording particle collisions at the LHC [2] at CERN. Being the most energetic particle accelerator, the LHC still offers unique physics opportunities for studying the Higgs and electroweak sectors of the standard model[3]; if the study at the energy frontier is to continue, the LHC must go on. The HL-LHC upgrade aims to increase the luminosity of the LHC by up to a factor of 7 in the next 10 years, which ultimately increases the number of meaningful collisions. Naturally, various sub-systems of the experiments used to capture the outcomes of the collisions will require upgrades to handle higher collision rates and background radiation rates than they were designed for.

The ATLAS experiment [4] is one of the LHC’s general-purpose particle detector arrays, positioned around one of the collision points of the LHC. It detects the products of LHC collisions. During the 2019-2022 Long Shutdown of the LHC, the most complex upgrade of the ATLAS experiment is the replacement of the small wheels of the muon spectrometer with the so-called New Small Wheels (NSWs) [5]. The NSW upgrade addresses both the expected decrease in hit efficiency in the precision tracking detectors of the current small wheel expected and the high fake trigger rate of the muon spectrometer. Two different detector technologies will be installed, stacked on the NSW frame: micromegas (MMs) and small-strip thin gap chambers (sTGCs). MMs are optimized for precision tracking while sTGCs are optimized for rapid triggering, although each will provide complete coverage and redundancy over the area of the NSW. Canada was responsible for providing 1/4 of the required sTGCs.

To reduce the fake trigger rate, the NSW will provide better track angular resolution to the ATLAS trigger system to reject tracks that do not originate from the collision [5]. sTGCs provide 100  $\mu\text{m}$  position resolution per detector plane [6], and are stacked in four (called an

sTGC quadruplet) to provide 1 mrad angular resolution on tracks [5, 7]. To be fast enough to provide this information to the trigger, they were designed with the smallest number of readout electrodes [5]. sTGCs are gas ionization chambers where a thin volume of gas is held between two cathode boards. One board is segmented into pad electrodes (of varying areas around  $300\text{ cm}^2$ ) and one is segmented into strip electrodes of 3.2 mm pitch. The signals picked up by the pads due to a passing charged particle are used to select which strip electrodes to readout. The position of the particle track in the precision coordinate can then be reconstructed from the strip signals to provide the track angle to the trigger fast enough for the trigger to decide if the particle came from the interaction or not [5].

Precise position resolution is gained without accurate positioning of readout electrodes in ATLAS. The ATLAS alignment system is able to position the surface of three sTGC or MM quadruplets traversable by a muon track with respect to one another within 40  $\mu\text{m}$ . The internal geometry of the detectors must be controlled or corrected for to within the chambers' position resolution [5]. Corrections to the position of strip electrodes in sTGC quadruplets are in their final stages. The corrections are done with characterization data collected throughout the construction process. At the cathode board level, strip electrode positions are digitized with a coordinate measuring machine (CMM) [8]. At the quadruplet level, sTGC quadruplets are characterized with cosmic rays and with an x-ray gun at positions that will be tracked by the alignment system. Cosmic muon data (cosmics data) can be used to measure relative strip position offsets in a chosen area with respect to the strip patterns on other layers, which characterizes the strips' alignment but does not allow the strips to be positioned in the absolute ATLAS alignment system. The x-ray method [9] is able to measure offsets of the strip pattern near the x-ray gun in a coordinate system accessible to the alignment system; however, it is limited to a handful of positions on the surface of the quadruplet and should be validated by an independent method. In this work, cosmics data is used to measure relative strip offsets, the x-ray method is validated with cosmics data, and how this work fits into the overall alignment scheme is presented.

Chapters 2 and 3 give more details on ATLAS, the LHC, the NSW, and sTGCs required to understand the context of this work. In chapter 4, the cosmic ray testing procedure and how the position of the strips can be probed with cosmics data is presented. Chapter 5 introduces the x-ray method, and in chapter 6, the x-ray offsets are validated with cosmic muon data. The thesis concludes with a summary and outlook in chapter 7.

# Chapter 2

## High energy particle physics at the LHC and ATLAS experiment

The LHC, ATLAS, and the upgrades they are undergoing are all motivated by the study of the Standard Model (SM) of particle physics and the open questions the SM does not address. Particle physics aims to study the indivisible constituents of matter. Understanding the fundamental building blocks and how they interact informs our understanding of the evolution of the Universe from the Big Bang to the forms of matter we recognize today – from atoms to the large scale structure of the universe. The reader is referred to [57, 56, 13] for more information on particle physics and the SM than is presented here. The LHC and ATLAS experiment are also introduced here to provide context for the NSW upgrade.

Do I need to place the background texts I cite in the in-text citations? I've listed where I've personally gotten each piece of information, but info about the SM and BSM might be so basic as to not require a citation. Comments?

### 2.1 The Standard Model

The SM describes all the fundamental particles and their interactions [GRIFFITHS]. It is a collection of quantum field theories that is able to explain the existence of all the particles discovered in the past century and predict how they interact to incredible precision [PESKIN]. In fact, when it was being developed in the 1960s-1970s, it began motivating the search for yet undiscovered particles with accelerators, like the Higgs Boson [10].

The standard model (represented in figure 2.1) consists of six quarks, six leptons, and five force-particles (and all their anti-particles). The force-carrying bosons are exchanged be-

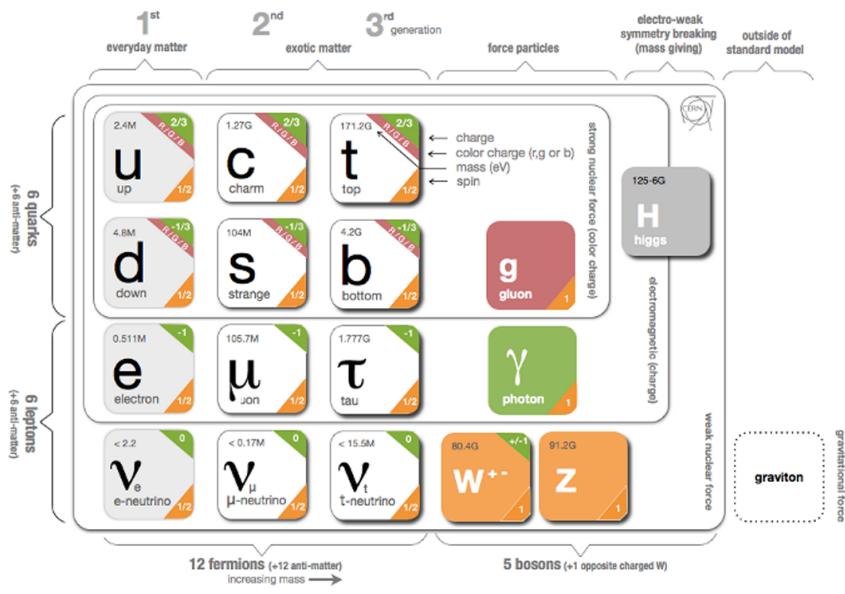


Figure 2.1: Representation of the standard model of particle physics, highlighting the three main categories of particles (quarks, leptons and force-particles) and the force-particles they interact with [11].

tween interacting particles to produce what is perceived as: the strong force mediated by gluons; the electromagnetic force mediated by photons; and the weak force mediated by the charged W-bosons and neutral Z-boson [GRIFFITHS]. The SM actually presents a theory of the electromagnetic and weak force as one force stemming from the same phenomena: the unified electroweak force [PESKIN]. The Higgs boson field interacts with the particles mediating the unified electroweak force in such a way as to distinguish the weak and electromagnetic forces from each other and give all particles (except neutrinos) a mass. This is referred to as electroweak symmetry breaking [PESKIN].

Quarks are fermions that are sensitive to all forces, but notably are the only particles sensitive to the strong force [GRIFFITHS]. Protons and neutrons are made up of quarks and gluons, and the strong force is responsible for their existence and mutual attraction into nuclei [NUCLEAR]. Leptons are particles not sensitive to the strong force. Charged leptons include the electron, which is responsible for the chemistry once part of atoms. Neutrinos are neutral, almost massless particles that only interact through the weak force [GRIFFITHS].

Common matter is made up of the lightest constituents of the standard model: up and down quarks, electrons and photons. The other particles are or were generated in high-energy environments and decay eventually to the lightest constituents. Such high energy environments include the Big Bang [BIG ORANGE BOOK], astrophysical sources, and accelerators [REV PART PHYS, GRIFFITHS]. The presence of the particles of the standard model at the beginning of the Universe means that their interactions and decays are fundamental for the study of the origin of the Universe [BIG ORANGE BOOK]. Many high energy astrophysical sources, like supernovae, generate particles that rain down on Earth as cosmic rays [34] [BOEZIO, REV PART PHYS, GRIFFITHS]. Finally, accelerators were built to create controlled high energy environments where the fundamental particles can be produced, detected and subsequently studied [GRIFFITHS, REV PART PHYS].

## 2.2 Beyond the Standard Model

Many open questions in particle physics remain that the SM provides no explanation for.

First, the fundamental force of gravity is not included in the standard model. One might expect a force particle to exist that mediates the gravitational force, but the weak strength of the gravitational force means the graviton will elude detection for a very long time. Moreover, there is no theory of gravity that does not require one of dark energy or dark matter [REV PART PHYS]. The universe is expanding at a rate irreconcilable with the known energy density of the Universe, and the nature of this "dark energy" is unknown [BIG ORANGE

BOOK]. Dark matter is the name given to mass whose effect on galaxies and our Universe is measurable through gravity, but for which there is no SM explanation [Munoz?].

Second, neutrinos in the standard model are massless; they do not interact with the Higgs field. However, in YEAR neutrino oscillations were measured, which can only occur if neutrinos do have mass [CITE].

Third, the unification of the electromagnetic and weak force begs the question of if there is a Grand Unified Theory that includes the strong force.

Theories beyond the standard model (BSM) aim to answer these questions. Often, BSM theories predict new particles. For example, super symmetry (SUSY) predicts that each SM particle has a heavier super symmetric partner. SUSY would explain the origin of dark matter with weakly interacting massive particles and BLANK. Ideally, a BSM theory predicts a measurable signature that can be search for at accelerators or elsewhere.

## 2.3 Studying high energy particle physics with accelerators

Accelerators of increasingly high energy have a long history of enabling the discovery of new particles [GRIFFITHS]. Only calling on one example, one of the main goals when the Large Hadron Collider and the ATLAS experiment were proposed was to detect the long-predicted Higgs boson particle – a triumph accomplished in YEAR. Being the last particle of the Standard Model to be discovered, the discovery marked the completion of the standard model as it is known today.

Since then, measurements of the cross section of particle physics processes have been enabled by the LHC that test our understanding of the standard model and provide hints of where to look for new physics. A summary of the cross sections measurements done at the ATLAS experiment is shown in figure 2.2.

Given the precision to which SM parameters like cross sections have been measured, any discrepancy between theory and experiment could be a portal to new physics tackling the biggest questions in particle physics today.

Accelerators and detectors can also be used to search for signatures of BSM theories. The controlled, high rate environment enables the search for rare processes that would be impossible to discern in other environments. If the signature is not found, exclusion limits can be set. [EXAMPLE OF FAMOUS ATLAS SEARCH].

Overall, accelerators play a key role in making precision measurements, searching for rare processes predicted by the SM, and testing BSM theories.

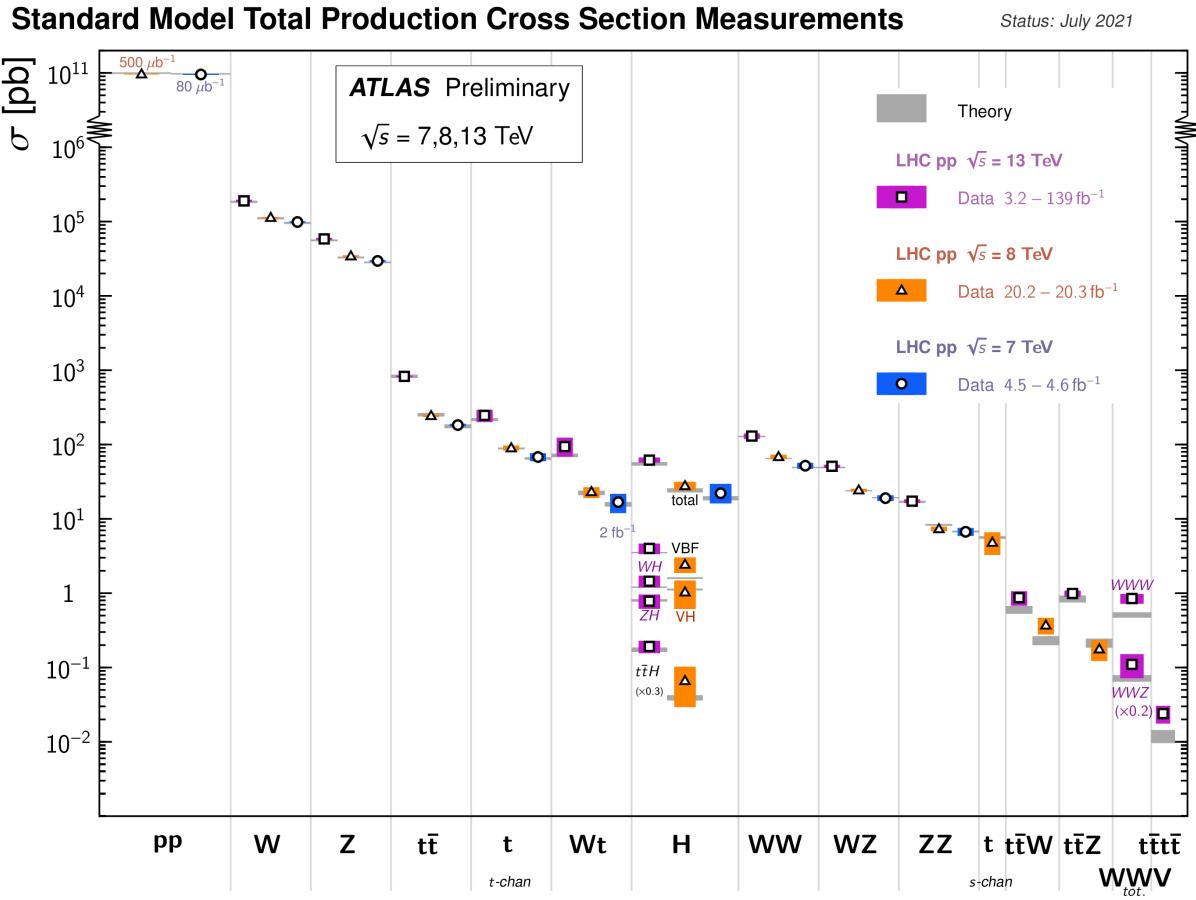


Figure 2.2: Cross sections of select SM physics interactions measured using the ATLAS experiment at the LHC. The comparison with theoretical predictions is also shown [12].

## 2.4 The Large Hadron Collider

The LHC is an accelerator 27 km in circumference and located  $\sim$ 100 m underground at CERN near Geneva, Switzerland [2]. It has two beam pipes that counter-circulate bunches of protons<sup>1</sup> before colliding the bunches in the center of one of four major experiments, such as the ATLAS experiment (discussed in section 2.5). There, the partons interact and many physics processes can occur. The LHC enables physicists to study particle physics at the energy frontier. In the previous run of the LHC (run-2), protons were collided with a center of mass energy of 13 TeV.

The number of proton-proton interactions generated by the LHC directly affects the statistics available to make fundamental measurements of cross sections, event rates, etc. Predicting the number of proton-proton interactions requires defining a metric called luminosity [13]. It is the number of particles an accelerator can send through a given area per unit time. It is calculated from the measurable quantities in equation 2.1.

$$\mathcal{L} = \frac{f N_1 N_2}{4\pi\sigma_x\sigma_y} \quad (2.1)$$

In equation 2.1,  $f$  is the frequency of the bunch crossings (25 ns),  $N_1$  and  $N_2$  are the number of protons in each bunch ( $\sim 10^{11}$  protons / bunch), and  $\sigma_x$  and  $\sigma_y$  are the RMS of the spatial distributions of the bunch. Therefore, luminosity is a property of the accelerator and its operating parameters. The design luminosity of the LHC was  $10^{34}$  cm $^{-2}$ s $^{-1}$ . The units of luminosity are an inverse area; multiplying the luminosity by the cross section of a given process gives the expected rate for that process.

Integrating the *instantaneous* luminosity (equation 2.1) over a period of data collection gives the integrated luminosity,

$$L = \int \mathcal{L}(t) dt \quad (2.2)$$

which is related to the total number of interactions. In this way, the luminosity is the link between the accelerator and the statistical power of measurements to be made with the data collected.

So far, the LHC provided an integrated luminosity of  $28.26 \text{ fb}^{-1}$  in run-1 [14] and  $156 \text{ fb}^{-1}$  in run-2 [15], as shown in figure 2.3. The HL-LHC upgrade [1] was accepted because without

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<sup>1</sup>the LHC also accelerates lead ions, but this paper is written in the context of proton-proton collisions



Figure 2.3: LHC/HL-LHC plan [17]. The integrated luminosity collected and projected for each run of the LHC is shown in red below the timeline and the center of mass energy of the collisions is shown in red above the timeline. The top blue arrow labels the run number. “LS” stands for “long shutdown” and indicates periods where the accelerator is not operating. During the shutdowns, upgrades to the LHC and the experiments are being installed. This timeline was last updated in January, 2021, and reflects changes in the schedule due to the ongoing pandemic.

increasing the luminosity of the LHC tenfold, running the accelerator will not provide significant statistical gain on measurements. Also, some systems will need repair and replacement to operate past  $\sim 2020$ . The LHC will be the most energetic accelerator in the world for years to come and is the only accelerator capable of directly producing the Higgs boson and top quarks, so the European Strategy for Particle Physics made it a priority “to fully exploit the physics potential of the LHC” with “a major luminosity upgrade” [16]. The goal is for the HL-LHC to provide an integrated luminosity of  $3000 \text{ fb}^{-1}$  in the 12 years following the upgrade. The luminosity actually achieved will depend on a combination of technological advances and upgrades in progress that affect the factors contributing to luminosity defined in equation 2.1 [1].

The increase in statistics the HL-LHC will improve measurements of standard model pa-

rameters and improve searches for unobserved phenomena [3]. A particular measurement that will benefit from the increased statistics is of the triple-Higgs coupling. Measuring the coupling measures the shape of the Higgs potential responsible for electroweak symmetry breaking. Any discrepancy with the SM prediction will show that there must be other sources of electroweak symmetry breaking – if there is such another source, it could be an avenue to study the other questions the SM does not address. Also, the sensitivity to several exotic Higgs decays to BSM particles will be improved. Many of the products of exotic Higgs decays are dark matter candidates or otherwise could extend the scope of the SM.

## 2.5 The ATLAS experiment

The ATLAS experiment [4] was designed to support all the physics goals of the LHC. It is 44 m long and 25 m in diameter, and weighs 7000 tones. It is an array of particle detector subsystems arranged concentrically around the beam pipe and centered around one of the LHC’s interaction points (a place where the beams collide), as shown in figure 2.4. ATLAS is cylindrical because it aims to provide  $4\pi$  coverage around the interaction point. It is helpful to separate the subsystems of ATLAS into the so-called “barrel” and “endcap” or “forward” regions.

For analysis, ATLAS is typically described in spherical coordinates. The azimuthal angle  $\phi$  is measured around the beampipe and the polar angle  $\theta$  is measured from the beam pipe. A more useful coordinate than  $\theta$  is the pseudo-rapidity,  $\eta = -\ln \tan(\theta/2)$ , because it approaches the rapidity of a particle when its momentum is much greater than its mass and differences in rapidity are approximately invariant to a Lorentz boost parallel to the beam. The range of  $\eta$  is 0 (perpendicular to the beam) to  $\pm\infty$  (parallel to the beam, or the z-direction). Typically,  $\eta$  is the physically interesting coordinate because the  $\phi$  coordinate follows the cylindrical symmetry of the beam.

It is not actually the proton bunches that collide and generate physics processes, but their constituent partons. Since the partons may carry an unknown fraction of the momentum, ATLAS analyses are based on the sum of the transverse momentum ( $p_T$ ) and energy of outgoing particle being approximately zero (transverse meaning perpendicular to the beam). **The goal of ATLAS is to reconstruct the transverse momentum of each collision product to understand what happened in each collision. -*j* Do I have this right?** A brief overview of the sub-systems of ATLAS is given, starting from the system closest to the beam and moving outwards. Each sub-system is responsible for a different group of collision products.

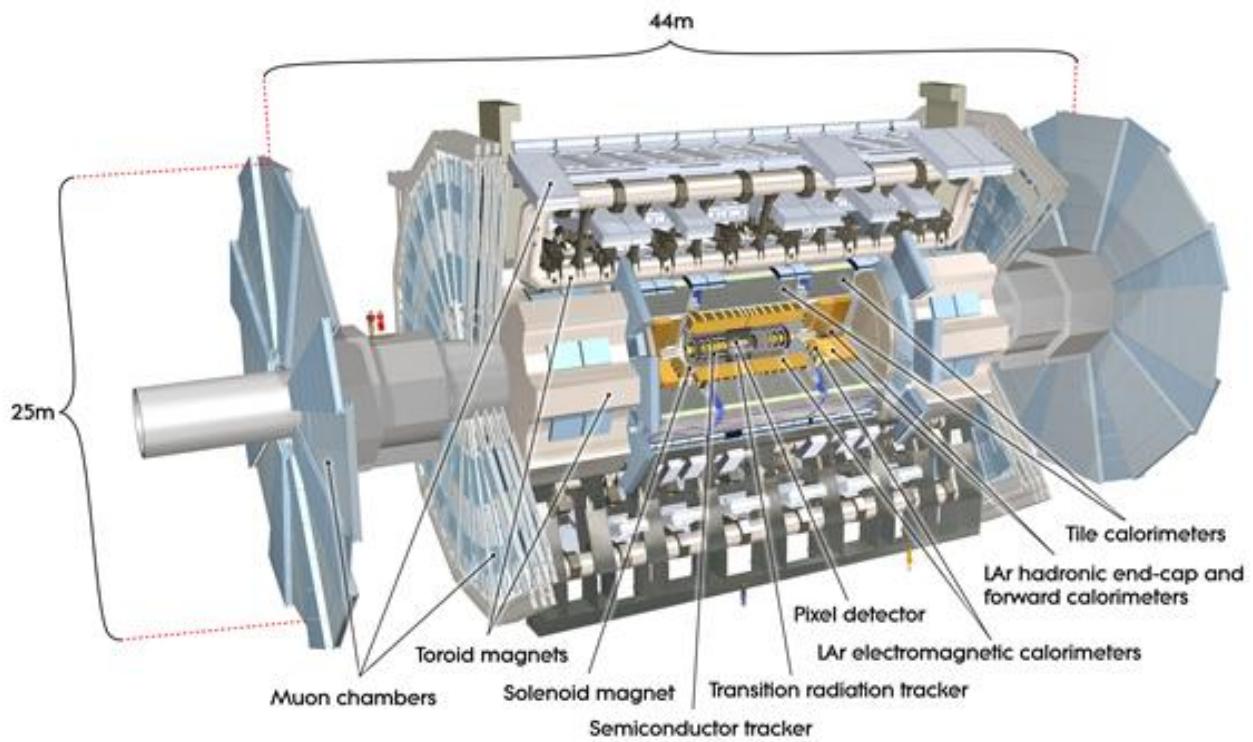


Figure 2.4: Diagram of the ATLAS experiment, with the various detector subsystems labelled. Figure from [4], which also contains more details about the ATLAS experiment.

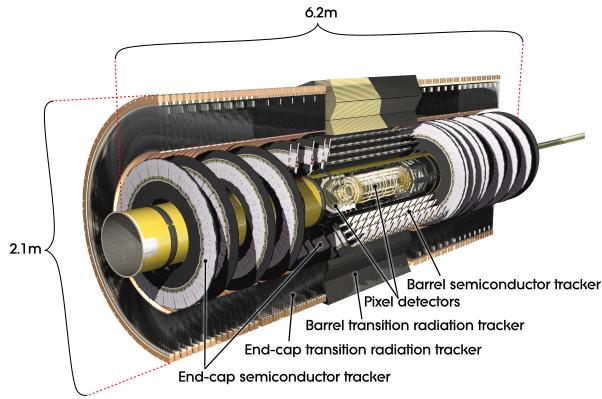


Figure 2.5: Diagram of the ATLAS experiment’s inner detector, with the different segments and the technology used labelled. Figure from [4].

### The inner detector

The inner detector [18, 19] (figure 2.5) is for precision tracking, vertex measurements and electron identification. A 2 T solenoid with field parallel to the beam bends the track of outgoing particles to make momentum measurements possible. The innermost part is made of high-resolution semiconductor pixel and strip detectors for precision tracking while the outermost part are straw-tubes that generate and detect transition radiation for electron identification.

### Calorimetry system

Electromagnetic and hadronic sampling calorimeter units are used to record the energy of electrons, photons, jets and missing transverse energy (from neutrinos, for example). A combination of liquid-argon (LAr) electromagnetic and hadronic calorimeters [20] and tile-scintillator hadronic calorimeters [21] cover the rapidity range  $|\eta| < 4.9$ , as shown in figure 2.6.

The calorimeters cause incoming charged particles to shower and deposit their energy in the sensitive volume. Only muons and neutrinos are known to pass the calorimeters to the muon spectrometer. Particles other than those mentioned would have decayed in the inner detector before reaching the calorimeter.

### Trigger system

It would be impossible to record all the data from bunch crossings every 25 ns, corresponding to a rate of  $\sim 40$  MHz, so ATLAS has a multi-level trigger system to select events of interest

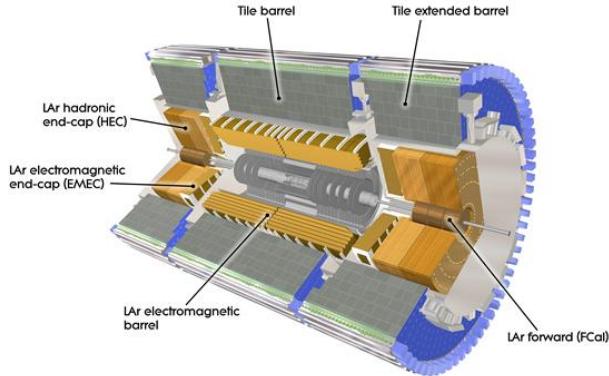


Figure 2.6: Diagram of the ATLAS calorimeter system, with the different segments and the technology used labelled. Figure from [4].

for permanent storage. The Level-1 (L1) hardware trigger [22] uses partial-granularity information from the muon spectrometer and calorimeter to trigger on high  $p_T$  muons, electrons, jets, high missing transverse energy, and  $\tau$  decaying to hadrons. The maximum L1 trigger rate ATLAS can accommodate is 100 kHz with a latency of 2.5  $\mu$ s.

The L1 trigger is used to define regions of interest that are fed into the software high level trigger (HLT), in which the full granularity of the muon spectrometer and calorimeter are used with information from the inner detector to reduce the trigger rate to 1 kHz. Events that pass the L1 and HLT trigger are recorded for use in offline analysis [23].

The ATLAS trigger system is described in the references above but the trigger rates quoted here are after the upgrades implemented for run-2, described in [24].

### Muon spectrometer

Magnetic deflection by superconducting air-core toroid magnets is used to measure muon momentum and energy. In the barrel of ATLAS, eight coils bent into “racetracks” are arranged around the beampipe provide the magnetic field. In the forward region, two end-cap toroids each with eight smaller racetrack-shaped coils arranged symmetrically around the beampipe are inserted in the ends of the barrel toroid [25]. Figure 2.7 shows the toroid magnets and the different parts of the ATLAS muon spectrometer.

The muon spectrometer [27] is separated into detectors used for precision offline tracking and for triggering. Three layers of monitored drift tubes (MDTs) or cathode strip chambers (CSCs) are used for tracking. The position of the muon track in each of the three layers allows reconstruction of the track sagitta and hence momentum. For the design momentum

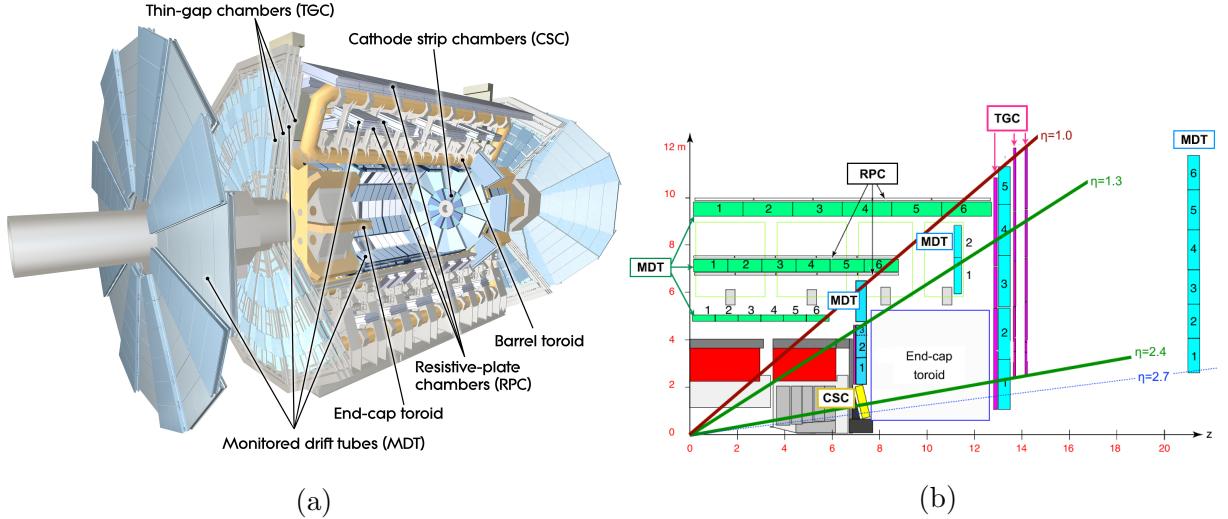


Figure 2.7: (a) The ATLAS muon spectrometer [4]. (b) A quarter-cut of ATLAS, with the interaction point in the bottom left corner. The small wheel is just left of the end cap toroid, the big wheel is to its right, and the outer wheel is the rightmost structure [26]

resolution of  $\Delta p_T/p_T < 1 \times 10^{-4}$   $p_T$  / GeV for  $p_T < 300$  GeV and a few percent for lower  $p_T$  muons, the MDTs and CSCs required position resolution of 50  $\mu\text{m}$  each. Accordingly, an optical alignment system was designed to monitor and correct for chamber positions [27, 28].

Resistive plate chambers (RPCs) are used for triggering in the barrel and thin-gap chambers (TGCs) are used for triggering in the endcaps. The positions of each type of chamber are sketched in figure 2.7b. Often, the endcap muon spectrometer is separated into three wheels – the small wheel (SW), big wheel, and outer wheel – ordered by proximity to the interaction point. In run-1, low (high)  $p_T$  muons were triggered on at L1 if two (three) of the RPCs or TGCs layers around the big wheel fired in coincidence, for the barrel and endcaps respectively [22]. After run-1 it was discovered that up to 90% of the triggers in the endcap were fake, caused by background particles generated in the material between the small wheel and the big wheel [5]. To reduce the fake rate in run-2, the TGCs on the inside of the small wheel also had to register a hit. The added condition reduced the trigger rate by 50% in the range  $1.3 < |\eta| < 1.9$  [24]. The effectiveness of the solution was limited since the  $|\eta|$ -range of the small wheel TGCs was limited to  $1.0 < |\eta| < 1.9$  and the position resolution of the small wheel TGCs is coarse [5].

# Chapter 3

## The New Small Wheels

### 3.1 Motivation for the New Small Wheels (NSWs)

The hit rate of all detector systems will increase with the HL-LHC not only because of the increase in luminosity, but also because the background radiation rate increases linearly with luminosity. The combined rate presents problems for both the tracking and triggering capabilities of the muon spectrometer [5].

In term of tracking, the efficiency of the MDTs decreases by 35% (mostly due to long dead-times) already when exposed to the maximum hit rate at the current luminosity, 300 kHz. At the threefold increase in luminosity predicted for run-3, most of the small wheel will be subjected to a hit rate well above 300 kHz and it will begin missing hits. Losing hits in the small wheel will reduce the high  $p_T$  muon momentum resolution. The decrease in resolution will affect the ability to search for, for example, high mass  $Z'$ ,  $W'$  and pseudo-scalar Higgs [5].

Already, the forward muon trigger system copes with a very high fake rate, even when including TGC data from the small wheel in the trigger as in run-2. At the luminosity expected in run 3, 60 kHz of the maximum 100 kHz of the L1 trigger would be taken by the endcap muon spectrometer. A possible solution would be to raise the minimum  $p_T$  threshold from 20 GeV to 40 GeV, but the ability to study several physics processes of interest depend on low  $p_T$  muons, particularly the Higgs to two muons decay and SUSY particle decays to leptons [5]. The ability to detect Higgs decays well is particularly important because that is how triple-Higgs events will be identified to measure the triple Higgs coupling [CITE].

The NSW will solve both these problems. It will be covered with precision tracking chambers suitable for the expected hit rates and triggering chambers capable of 1 mrad angular resolution. The idea behind the triggering chambers is to match the small wheel track segment

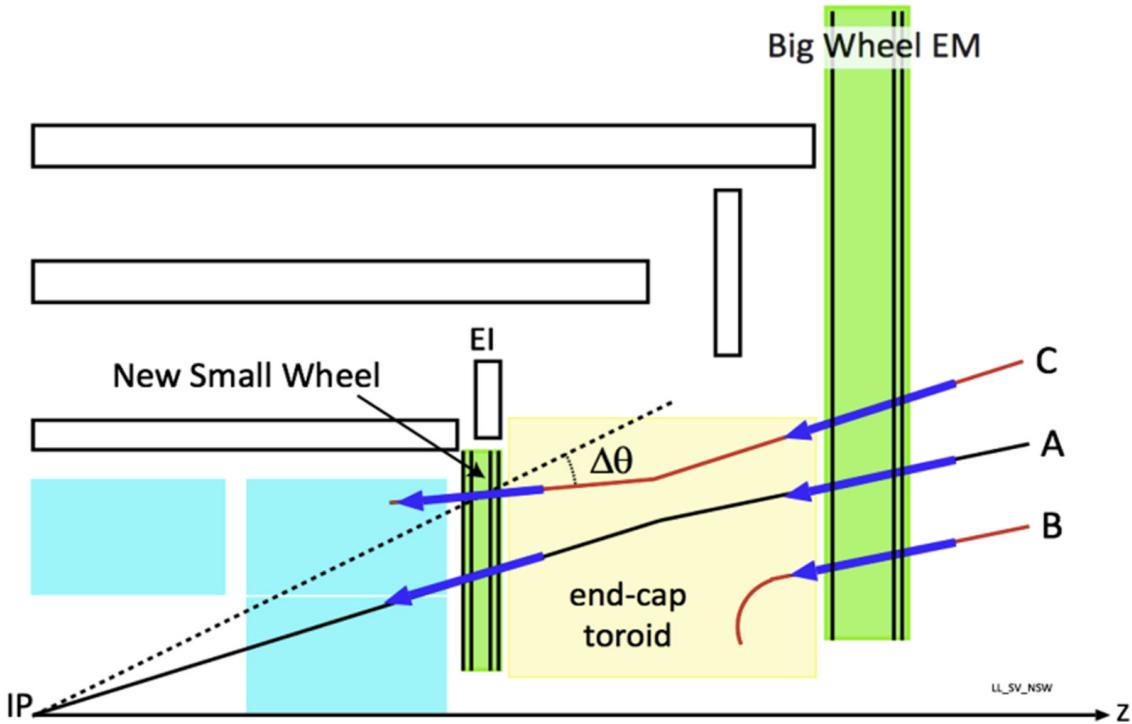


Figure 3.1: A schematic of a quarter cross section of the ATLAS detector, with the collision/interaction point (IP) in the bottom left corner. Three possible tracks are labelled. Ideally, track A would be triggered upon while track B and C discarded. With the small wheel, all three tracks would be recorded. With the NSW, only track A would be recorded [5].

with the track segment from the big wheel to discard tracks not originating from the interaction point. Figure 3.1 illustrates this point: the run-1 trigger system would have triggered on all three tracks, while with the NSW the trigger system would only trigger on track A [5].

## 3.2 Design of the NSWs

The NSWs are covered with two detector technologies: micromegas (MM) and small-strip thin gap chambers (sTGCs). MMs are the primary tracking detectors and sTGCs are the primary triggering detectors, but for redundancy sake both are designed to do either. As such, both sets of detectors are to have position resolution better than  $\sim 100 \mu\text{m}$  per plane. Four chambers of each type are glued together to create quadruplet modules. Quadruplets of different sizes are assembled into wedges. Two sTGC wedges and two MM wedges are

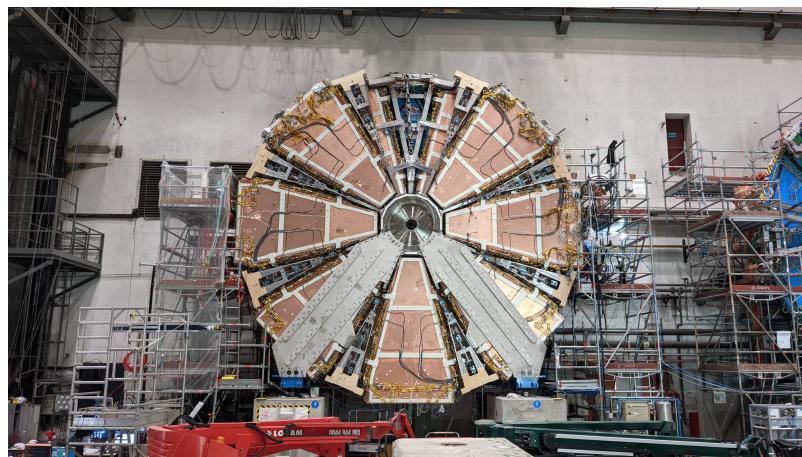
layered to create sectors (with the sTGC wedges on the outside) [5]. Different stages of the constrcution process are shown in figure 3.2.



(a) An sTGC quadruplet module. The left image highlights the trapezoidal shape. The right image shows the short edge corner. The four sTGC layers and each layer's gas inlet are visible. The gas outlets and high voltage cables are at the long edge in the back of the photo. The green printed circuit board along the sides are the adaptor boards where the front end electronics are attached.



(b) Left: An sTGC wedge. The white frame outlines the individual quadruplet modules. Right: A completed sector, with two sTGC wedges on the outside and two MM wedges on the inside.



(c) The new small wheel. All sectors except one large sector at the top are installed, revealing two of the smaller sectors that are normally hidden under the large sectors and support bars. The NSWs are 10 m in diameter.

Figure 3.2: Images breaking down some of the construction units of the NSWs.

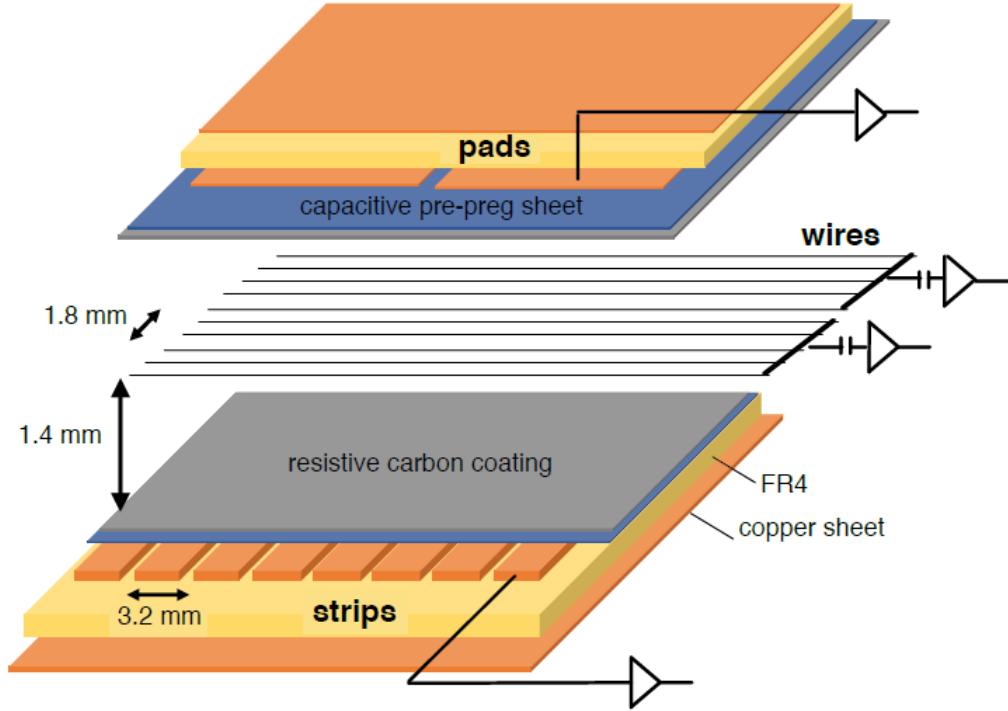


Figure 3.3: Internal structure of an sTGC, zoomed into the area under approximately 1 pad. High voltage is applied to wires suspended in the gas volume to create an electric field. A passing muon causes an ionization avalanche that is picked up by the wire, strip and pad electrodes [9].

### 3.3 Small-strip thin gap chambers

sTGCs are gas ionization chambers operated with a CO<sub>2</sub>:n-pentane ratio of 55:45. Gold-plated tungsten wires, 50 µm in diameter and with 1.8 mm pitch, are suspended between two cathode planes made of FR-4, each 1.4 mm away (see figure 3.3). One cathode board is segmented into pads of varying area (around 300 cm<sup>2</sup> each), and the other segmented into strips of 3.2 mm pitch, perpendicular to the wires. High voltage is applied to the wires and the cathode planes are grounded [5, 7]. When a muon passes through, the gas is ionized and the electric field in the gas gap causes an ionization avalanche [29]. The motion of the ions and electrons are picked up on the nearby wire, strip and pad electrodes. The resistivity of the carbon coating and capacitance of the pre-preg sheet tune the spread of the charge

distribution [30] and the speed of the response [31] to optimize the rate capability. The hatching of the strips and wires establishes a coordinate system from which to extract the coordinate of the muon as it passes through the layer [5].

A 3-out-of-4 coincidence in the pad electrodes of a quadruplet will define a region of interest where the strip and wire electrodes should be readout. Then, a track segment with the required angular resolution can be constructed quickly enough to be provided as input to the hardware trigger [5]. The pad-triggering scheme greatly reduces the number of electrodes that require readout to provide the design track angular resolution of 1 mrad. For the next run of ATLAS (run-3), the L1 trigger will pass if the track segment from the small wheel matches a region of interest defined by the TGCs of the big wheel [32, 5]. The increased angular resolution and  $\eta$  coverage of the quadruplets will be an improvement in the triggering scheme over the current small wheel TGCs [5]. After the phase-II upgrade of the trigger and data acquisition system scheduled for  $\sim$ 2025, which includes upgrades to the big wheel’s angular resolution, there will only be a trigger if the combined track points back to the interaction point, increasing the muon  $p_T$  resolution [5, 33].

Signal is readout from groups of successive wires, so the position resolution in the direction perpendicular to the wires is 10 mm. The resolution in this direction is sufficient since it will give the symmetric azimuthal coordinate in ATLAS. Good resolution on the  $\eta$  coordinate, perpendicular to the strips, is important [5]. The average single chamber position resolution in the strip coordinate was 45  $\mu\text{m}$  for perpendicular muon tracks as measured in a test beam [6] — well within design specifications. When four sTGCs are glued together into a quadruplet the design angular resolution of 1 mrad in the strip coordinate is achievable [5, 7].

Therefore, sTGCs are able to meet the precision tracking and triggering goals they were designed for. To ensure they can deliver once installed in ATLAS, knowing the position of the strips to within their position resolution in the ATLAS coordinate system is necessary. The NSW alignment system, detailed in section 3.5 monitors the position of alignment platforms installed on the surface of the wedges. The alignment platforms are installed with respect to an external reference on the sTGCs: two brass inserts on each strip layer on one of the angled sides of each quadruplet (shown in figure 3.4). So the challenge of positioning the strips in ATLAS was separated into two steps: first, position the strips with respect to the brass inserts; second use the alignment system to position the alignment platforms. The next section provides some pertinent details on the sTGC construction process, with steps that affect the position of the strips with respect to the brass inserts highlighted.



(a) Brass insert near long edge.

(b) Brass insert near short edge.

Figure 3.4: The brass inserts sticking out from the gas volumes of an sTGC quadruplet. These inserts were pressed against alignment pins when the individual sTGCs were being glued together.

### 3.4 sTGC Quadruplet Construction

Five countries were responsible for producing the sTGC modules of varying geometries for the NSW: Canada, Chile, China, Israel and Russia. Canada was responsible for 1/4 of the required sTGCs, of three different quadruplet geometries. The steps of the construction process in each country were similar [5]. The process followed in Canada is detailed.

TRIUMF in Vancouver, British Columbia was responsible for preparing the cathode boards. The boards were made and the electrodes etched on at a commercial laboratory, Triangle Labs, in Carson City, Nevada. Once completed they were sent to TRIUMF to be sprayed with graphite and otherwise prepared [8]. The boards are commercial multilayer printed circuit boards, but the strip boards required precision machining to etch the strip pattern [5]. Triangle Labs also machined the two brass inserts into each strip board. A coordinate measuring machine (CMM) was used to digitize a set of reference strips. Four quality parameters describing non-conformities in the strip pattern of each board with respect to the brass inserts were derived from the data and the results are available on a QA/QC database. The parameters and the CMM data collection is described in full in Carlson's thesis [8]. Due to time constraints, tolerances on the non-conformities in the etched strip pattern with respect to the brass inserts were loosened, with the condition that the strip positions would have to be corrected for [8].

The prepared boards were sent to Carleton University in Ottawa, Ontario for construction into sTGCs and quadruplets. First, the wires were wound around the pad cathode boards using a rotating table and the wires were soldered into place. A wound pad cathode board was held by vacuum on a granite table, flat to within  $20\text{ }\mu\text{m}$  and a strip cathode board glued on top to create an sTGC. Holding one sTGC flat with the vacuum, another was glued on top to create a doublet, then two doublets were glued together to create a quadruplet. When gluing sTGCs together, the brass inserts were pushed against alignment pins with the goal of keeping the strip layers aligned within tolerance. However, non-conformities in the shape of the brass inserts, non-conformities in the position of the alignment pins and shifts between strip layers while the glue cured resulted in misalignments between the brass inserts – and the strips – on successive layers. Precise alignment of the pad boards or wires with respect to the strip boards did not have to be so tightly controlled because pads do not measure the precision coordinate.

The Carleton team finished the quadruplets by installing adaptor boards on the angled sides of each layer that allow front end electronics to be attached. Completed quadruplets were sent to McGill University where they were characterized with cosmic rays. The details of cosmic ray testing are described in chapter 4. Quadruplets were then sent to CERN where they were assembled into wedges and alignment platforms installed. The alignment platforms

were installed using a jig positioned with respect to the brass inserts. Completed wedges were assembled into sectors then installed on the NSWs.

The quadruplet construction process had two steps where strip positions could be shifted off of nominal. At board-level, there could be non-conformities in the etched strip pattern with respect to the brass inserts, described by the four quality parameters [8]. At the quadruplet level, misalignments between the brass inserts and strips on different layers were introduced during the gluing. The result was that the brass inserts were not a reliable reference point and that the strips can be offset from their design position by up to hundreds of micrometers. Offsets in strip positions from nominal in Canadian quadruplets were shown to be random [8], so no one correction would suffice. The offsets must be measured and corrected for in the ATLAS offline software, **Athena**, which does the precision tracking. Understanding the work ongoing to make measurements of offsets and correct for them requires understanding the strategy of the NSW alignment system.

## 3.5 NSW alignment

What figures can I use to show the bars, the alignment jig, the platforms and the optical fibres?

The idea of the NSW alignment system is presented in [5], but the details have only been presented internally so far. After the wedges are constructed, alignment platforms are installed on every sTGC quadruplet and optical fibres routed to them, as shown in figure 3.5. Light from the optical fibres will be monitored in real time by cameras (BCAMs) mounted on the alignment bars of the NSWs. The system will thus record the positions of the alignment platforms in the ATLAS coordinate system, accessible at any point during operation.

The original alignment scheme was to use the brass inserts as a reference between the alignment platforms and the individual strips, as shown in the solid arrows in figure 3.6 – this will no longer work. The position of the alignment platforms will be known thanks to the alignment system, so a different method to get the position of the individual strips with respect to the alignment platform is currently in its final stages. It uses the yet-unmentioned x-ray dataset to calculate offsets of the strip pattern of an sTGC layer in a local area (local offset) with respect to the nominal geometry by analyzing the beam profile left by an x-ray gun attached to different positions on the alignment platforms. The alignment platforms provide the link to the nominal geometry because their position with respect to the strips is known in the case that the strips are perfectly etched and aligned. Effectively, the reference to the brass inserts is skipped, represented as the dashed line in figure 3.6.

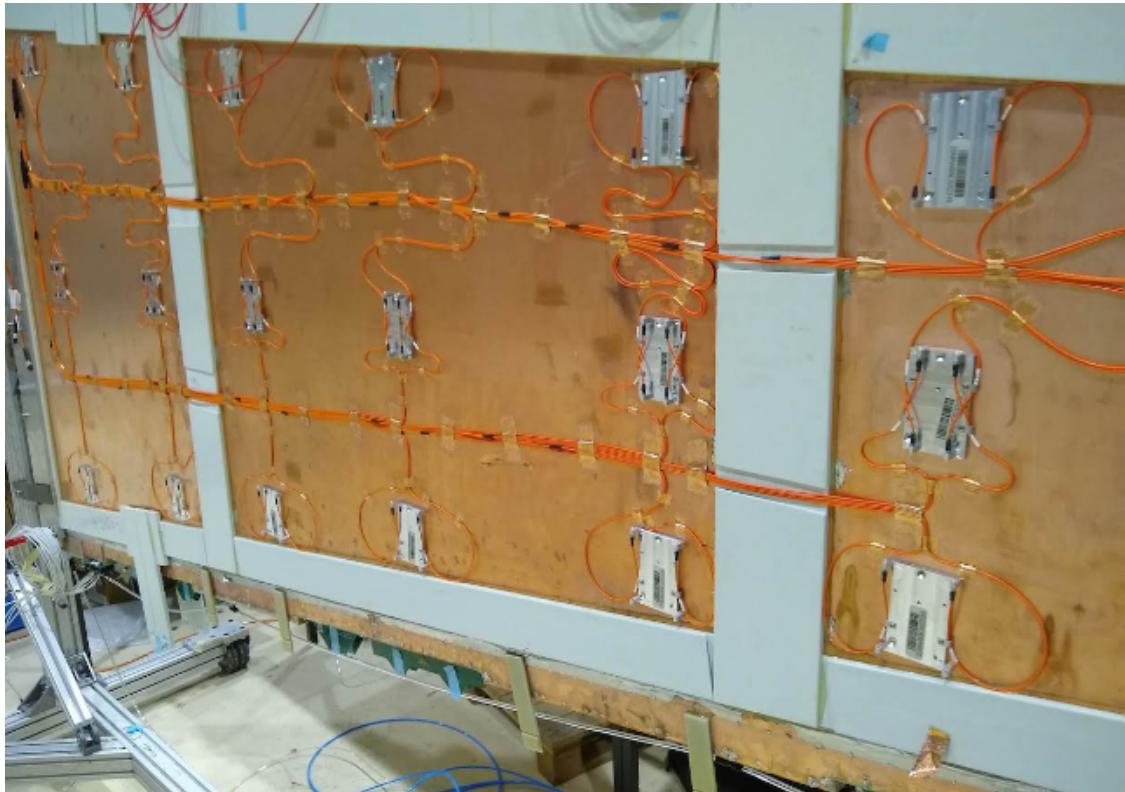


Figure 3.5: An sTGC wedge with alignment platforms (silver) installed on the quadruplet. Optical fibres (orange) are routed to the alignment platforms. Cameras on the frame of the NSW will record light from the optical fibres to position the alignment platforms in the ATLAS coordinate system.

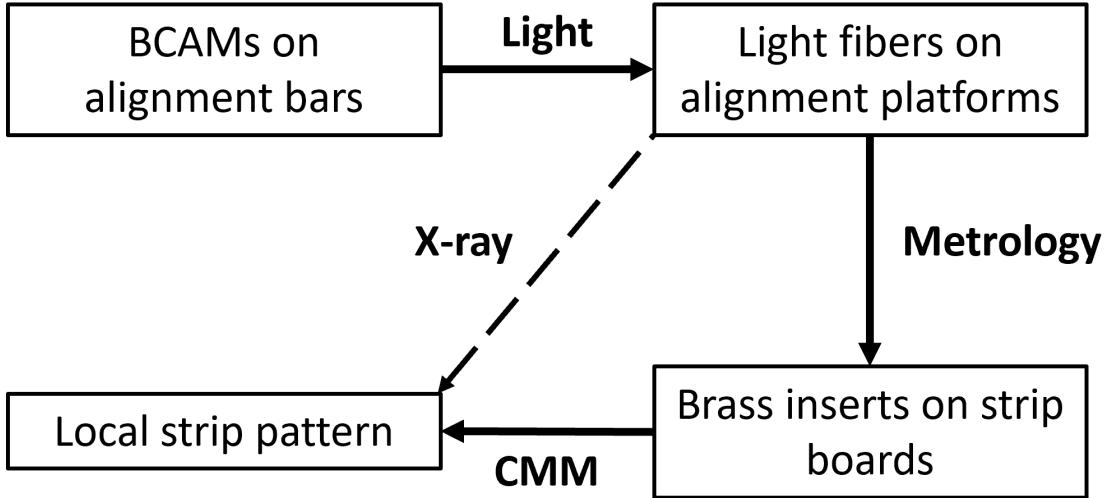


Figure 3.6: How the different elements of the sTGC alignment system relate to one another. The solid arrows denote the planned alignment scheme. The dashed arrow shows the modification being finalized now. This figure was originally designed by Dr. Benoit Lefebvre.

The x-ray method does not have the sensitivity to measure the offset of each strip from nominal, but instead in a local area around the position of the gun. Local offsets are thus used to build an alignment model for each strip layer. Formally defined, an alignment model is a set of parameters used to estimate the true position of a strip given its nominal position. The alignment model currently being worked on takes x-ray and CMM data as input to calculate a global offsets and rotations of each strip layer with respect to nominal [9]. Without the x-ray dataset, there would be no input to the alignment model that takes into account inter-layer misalignments introduced in construction.

Given that the x-ray local offsets can only be measured at positions where the gun can be attached and that they are an important part of the alignment scheme, the x-ray method needs to be validated. The goal of this thesis is to validate the x-ray local offsets while exploring how cosmics data complements and adds to the understanding of strip positions and overall alignment.

# Chapter 4

## Using cosmic muons to measure relative strip position offsets

At McGill, among other quality and functionality tests, each Canadian-made quadruplet was characterized with cosmic muons. In this chapter, the experimental setup and how the data was analyzed to provide relative strip position offsets is presented. The analysis method was motivated by the how it could be compared to data collected with the x-ray method (chapter 5) but also stands alone as a characterization of the alignment between strips of different layers. First, a brief introduction to cosmic ray physics.

### 4.1 Cosmic rays

The Earth is being bombarded by particles from the sun, galactic sources and extra galactic sources, collectively called cosmic rays [34, 13]. Cosmic rays are mostly protons, but also heavier ions, gamma rays and the term sometimes includes neutrinos. The primary (initial) cosmic ray interacts with the atmosphere causes electromagnetic showers and hadronic showers of particles. Hadronic showers result from the primary cosmic ray interacting strongly with the target of the atmosphere; the most abundant products are pions. Charge pions mostly decay to muons (there is a lesser contribution to the muon flux from kaons as well) [35]. Thanks to time dilation extending the muon's lifetime as measured on Earth, a flux of approximately 1 muon/cm<sup>2</sup>/ min reaches the ground [13]. Measuring the muon flux and energy spectrum reveals information about primary cosmic rays [35] which is interesting to high energy physicists and astrophysicists. The muon flux is also terribly convenient for testing muon detectors.

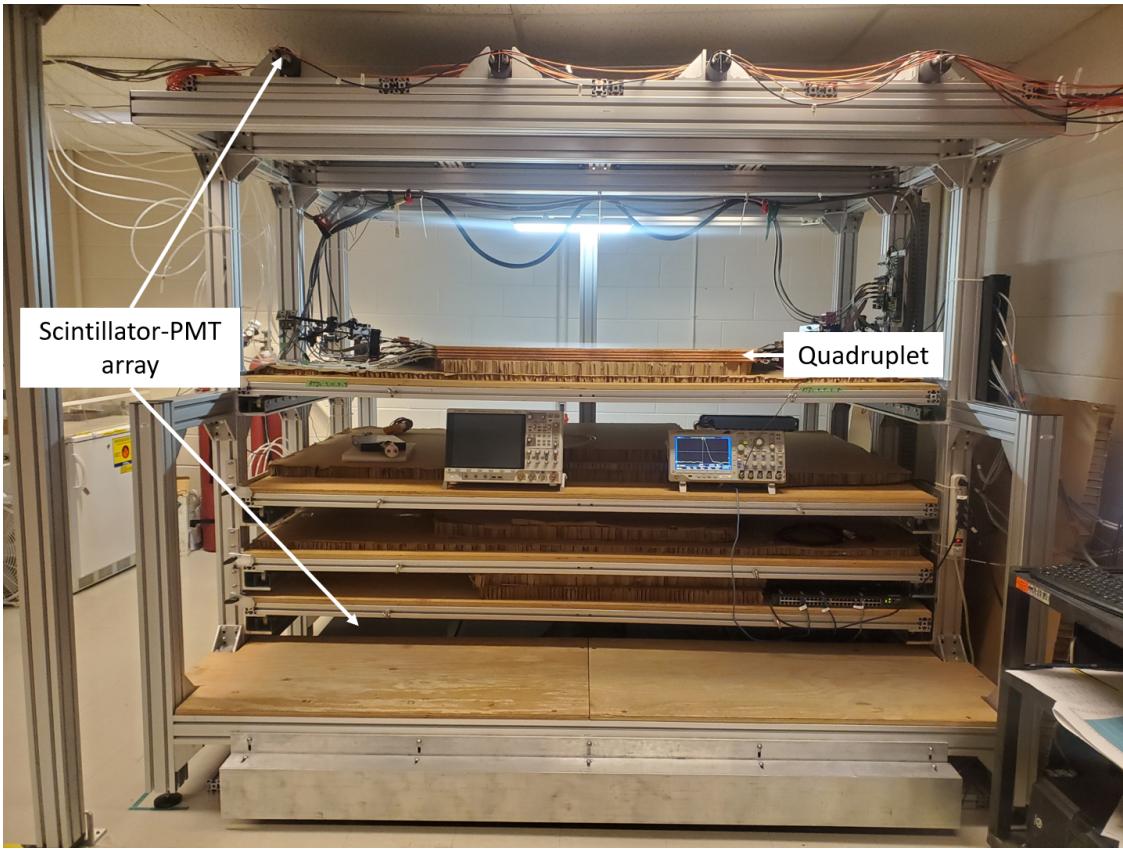


Figure 4.1: Cosmic muon hodoscope at McGill University with sTGC quadruplet in the test bench.

## 4.2 Experimental setup

Cosmic muon characterization was done with a hodoscope, a complete description of which can be found in [36]. The quadruplet was placed in the center of the test bench. Above and below it was a layer of scintillator-PMT arrays, labeled in figure 4.1. When a cosmic muon passed within the acceptance of the hodoscope, at least one scintillator from the top array and at least one from the bottom array fired in coincidence. The coincident signal was used to trigger the readout of the quadruplet's electrodes. The trigger was passed to the front end boards (FEBs) attached to the adaptor boards of each layer of the quadruplet.

Operating the chambers also required gas and high voltage. A pentane-CO<sub>2</sub> mixture was mixed and delivered to each sTGC with a gas system designed and made at McGill University. The gas system was controlled by a slow control program, also made in-laboratory [37]. To

prepare the quadruplets for operation, CO<sub>2</sub> was flushed through them overnight. Then, five gas volumes of the pentane-CO<sub>2</sub> mixture was flushed through (approximately 3 hours). High voltage was provided by CAEN boards.

## 4.3 Data acquisition

Each sTGC electrode was connected to a channel on a prototype ASIC<sup>1</sup> on the FEB, attached to the adaptor boards on each layer of a quadruplet. The ASIC amplified the signal was set to measure and record the signal peak amplitude from electrodes. For each trigger, the signal peak amplitude of all channels above threshold were recorded as an event and stored in a binary file. Thresholds were estimated [39] and adjusted manually in the configuration/readout software before the start of data acquisition. There was an exception to the threshold rule: the signals on strips adjacent to a strip above threshold were also readout using the so-called “neighbour triggering” function of the ASIC.

The quadruplets were held at 3.1 kV for approximately two hours to collect data from 1 million muon triggers.

## 4.4 Data preparation

### 4.4.1 Cuts on electrode hits

Corrupted data is removed while the raw data is being recorded in a binary file. The binary file is decoded into a usable ROOT [40] tree offline.

A hit is defined as a signal recorded from a channel that was above threshold or (in the case of strips) neighbour triggered. In addition to hits from muons, the quadruplets record noise from the electronics and  $\delta$ -rays. Therefore, cuts are applied to reduce the number of noise hits. The edge strips are very noisy, so all strip hits on layers with strip hits on either edge channel are cut. A default pedestal value is subtracted from the recorded signal peak amplitude of each electrode for a more realistic estimate of the signal amplitude. Also, events that only have hits on pad electrodes (no strips or wires) were cut because the large area of the pads made them susceptible to noise.

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<sup>1</sup>the VMM3 [38], designed for the MMs and sTGCs of the NSW

#### 4.4.2 Clustering and tracking

Many of the high-level characterization metrics required rebuilding muon tracks. For events passing quality cuts, the  $x$ - and  $y$ -coordinates of the ionization avalanche on each layer were extracted from the signal on the wires and strips respectively for each event, as is sketched in figure 4.2. In this work,  $x$  was the coordinate perpendicular to the wires and  $y$  was the coordinate perpendicular to the strips, as is drawn in figure 4.2.

The  $x$ -coordinate was taken as the center of the wire group with the maximum peak signal amplitude, since the wire groups' pitch (36 mm) was larger than the typical charge spreading. Assuming that the true  $x$ -position of the hit is sampled from a uniform distribution over the width of the wire group, the uncertainty in the  $x$ -position was given by  $\frac{36}{\sqrt{12}}$  mm = 10 mm [41].

The  $y$ -coordinate was taken as the Gaussian mean of the peak signal amplitude distribution across groups of contiguous strips. The process of grouping contiguous strip hits on a layer recorded in an event was called clustering, and the resulting group was called a cluster. Figure 4.2 sketches the clustering process and a sample cluster is shown in figure 4.3. The data acquisition system recorded the electrode ID of the strip hit, and in the clustering process the position of the center of the strip electrode is calculated based on the nominal quadruplet geometry. Typically, clusters were built of 3-5 strips. The thickness of the graphite coating over the cathode boards determined how many strips picked up the ionization image charge. Larger clusters were more likely caused by  $\delta$ -rays since they spread the cloud of ionization.

Events are cut from the analysis if there are two clusters on one layer's set of strips (indicative of noise). Clusters are cut if the cluster size is lesser than three (which should not happen for real events thanks to neighbour triggering), and if the cluster size is greater than 25.

The uncertainty in the  $y$ -coordinate could have been taken as the fitted cluster mean's statistical uncertainty; however, after comparing the difference in cluster means for different fitting algorithms in appendix A.2, 60  $\mu\text{m}$  of uncertainty was assigned.

The coordinates of the avalanches' on all layers were used to reconstruct tracks in  $x$  and  $y$  respectively. Most tracks passing quality cuts were probably cosmic muons, but of course some  $\delta$ -rays and noise contributed. The tracks were then used to calculate characterization metrics like electrode efficiency and spatial resolution, the details of which are discussed in [36].

### 4.5 Measuring relative local offsets

The offset of a strip from its nominal position can be modeled as a passive transformation. For each area of a strip layer, the local offset is the shift of the strip pattern in that area

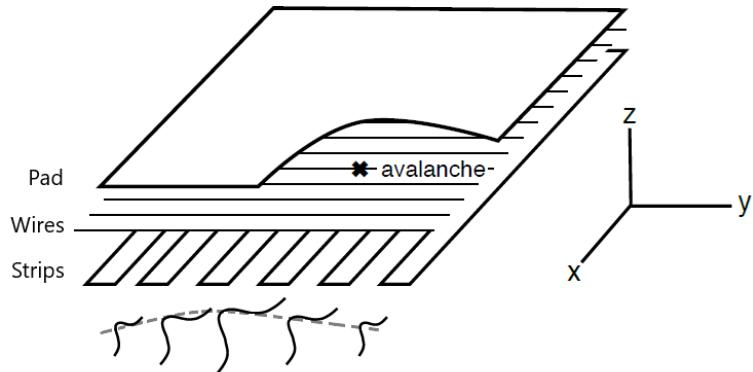


Figure 4.2: A sketch of an sTGC-like detector. The position of the avalanche could be extracted from the wires and strips that picked up the avalanche signal. The signals on individual strips are sketched. Clustering was the processs of fitting a Gaussian to the peak value of the signals on individual contiguous strips, as is done in figure 4.3. In this work, the  $x(y)$ -coordinate will always refer to the coordinate perpendicular to the wires (strips) [36, 30].

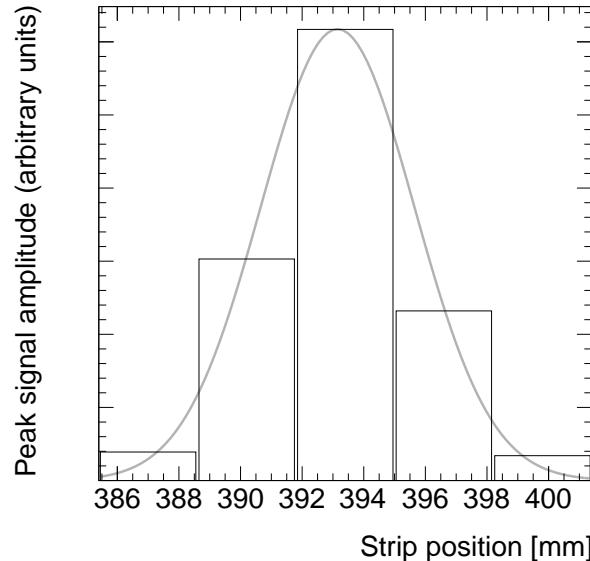


Figure 4.3: A sample cluster resulting from the current picked up on a group of strips after the passing of a muon (presumably). The grey line is a Gaussian fit.

with respect to the nominal geometry. Local offsets systematically change the position on the strips nearest to a passing muon. The `tgc_analysis/CosmicsAnalysis` software assumes that strips are in their nominal positions, so the recorded muon y-position ( $y$ ) is shifted opposite to the local offset ( $d_{local}$ ),

$$y = y_{nom} - d_{local}, \quad (4.1)$$

where  $y_{nom}$  is the position of the muon that would have been recorded if there was no local offset. Equation 4.1 ignores other factors that could affect the cluster position (like position resolution). With cosmics data, the local offset is unknown and there was no external reference to measure  $y_{nom}$ . Therefore, only relative local offsets could be calculated.

The minimal relative coordinate system uses two reference or fixed layers [36]. The hits on the two fixed layers were used to create tracks that can be interpolated or extrapolated (polated) to the other two layers. The set of two fixed layers and the layer of interest are referred to as a tracking combination. The residual of track  $i$ ,  $\Delta_i$  is defined as,

$$\Delta_i = y_{i,hit} - y_{i,track}, \quad (4.2)$$

where  $y_{i,hit}$  is the recorded hit position and  $y_{i,track}$  is the polated track position built from hits on the two reference layers. Track residuals are affected by the relative local offset in the area of each layer's hit. As an example, in figure 4.4, the residual on layer 2 perhaps indicates that layer 2 is offset with respect to layers 1 and 4 in the area of the track. Of course, a single track residual says nothing of the real relative local offset because of the limited spatial resolution of the detectors and fake tracks caused by noise or delta rays. However, the mean of residuals for all tracks in a region will be shifted systematically by the local offsets between layers [36]. For a quadruplet with nominal geometry, the mean of residuals should be zero in all regions and for all reference frames, unlike the example regions in figure 4.5. The value of the mean of residuals is a measure of the relative local offset of the layer with respect to the two fixed layers.

The residual distributions were wider for tracking combinations where the extrapolation lever arm was largest, as in the example distributions shown in figure 4.5. In general, residual means from distributions of residuals with geometrically less favourable tracking combinations have larger statistical and systematic uncertainties. The bin size of 200  $\mu\text{m}$  for the distributions shown in figure 4.5 was chosen based on the uncertainty on residuals calculated from tracks on layer 4 (1) built from hits on layers 1 and 2 (3 and 4) given a cluster y-position uncertainty of 60  $\mu\text{m}$  (appendix A.3), since these tracks yield residuals with the largest uncertainties.

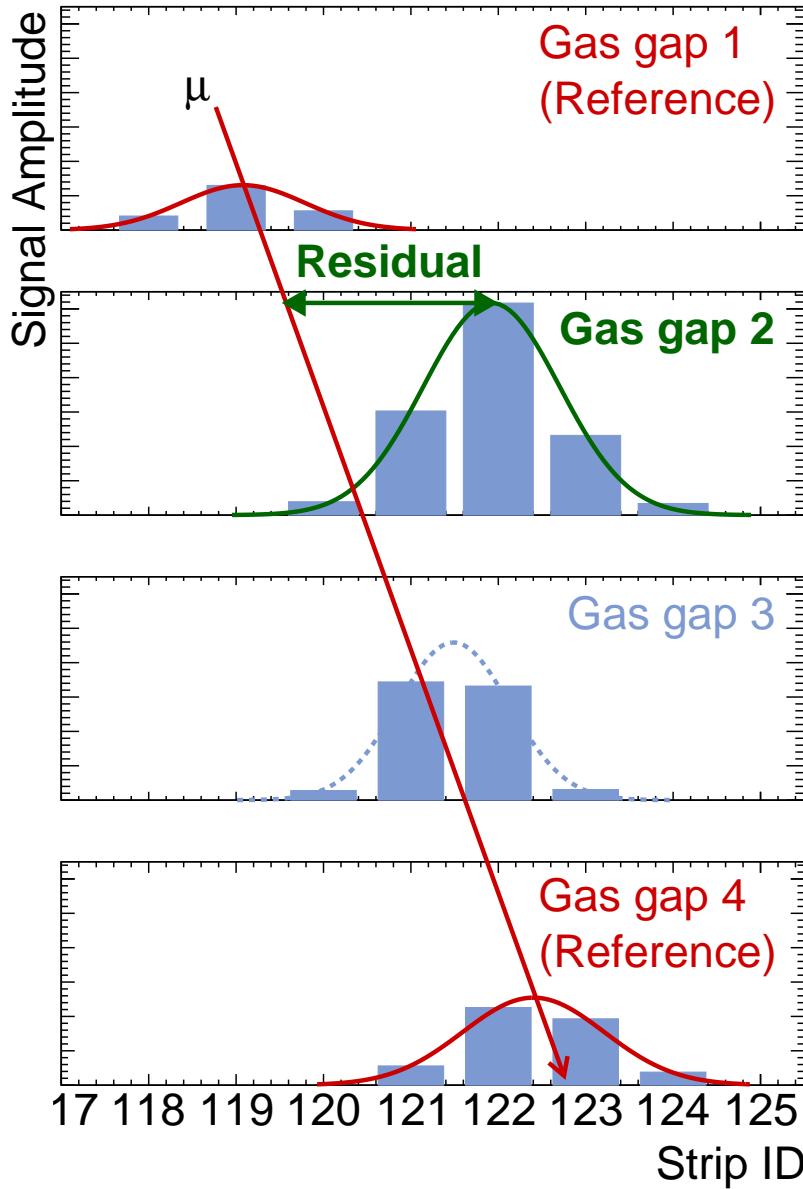
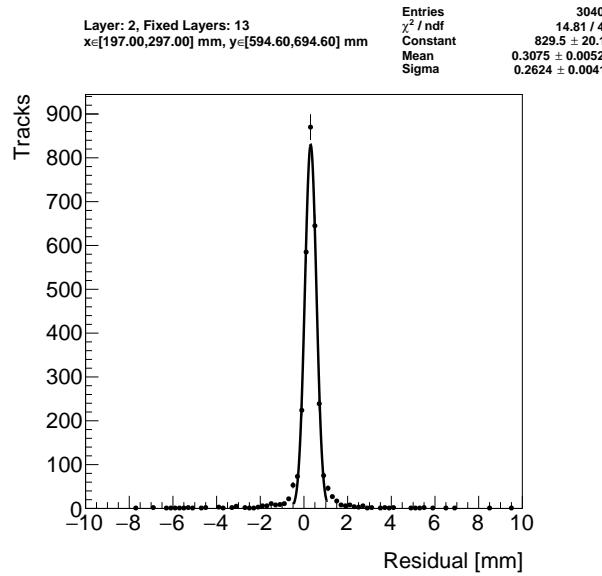
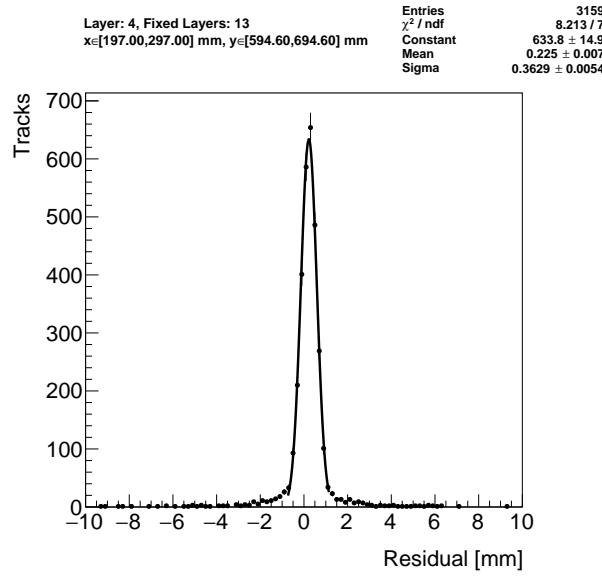


Figure 4.4: Representation of a muon event recorded by an sTGC. The clusters are fit with a Gaussian and the mean is taken as the hit position. A track is built from the chosen reference layers, 1 and 4, and the residual calculated on layer 2. The clusters come from a real muon event, but their positions were modified to highlight the residual on layer 2.



(a) Tracks on layer 2, reference layers 1 and 3.



(b) Tracks on layer 4, reference layers 1 and 3.

Figure 4.5: Residual distribution in the region  $x \in [197, 297]$ ,  $y \in [594.6, 694.6]$  mm (100 mm by 100 mm area) for two different tracking combinations.

A gaussian fit was used to extract the mean of the residual distributions. Theoretically, a double gaussian distribution is more apt, but for this analysis the gaussian fit was sufficient, as discussed in appendix C.1.

The area of the region of interest was 100 mm by 100 mm. The size balanced the amount of tracks falling in the region of interest to give sufficient statistics to the local residual distributions, while being smaller than the order on which local offsets were expected to change significantly. The change in local offsets over the surface of a layer can be modelled using global alignment parameters. Using the base alignment model with a global offset and rotation of each strip layer, “significantly” was defined based on the distance in  $x$  that a large but possible rotation of 1000  $\mu\text{rad}$  would change the local offset by more than 50  $\mu\text{m}$  – half the required position resolution of the sTGCs [5].

## 4.6 Visualizing relative alignment between layers

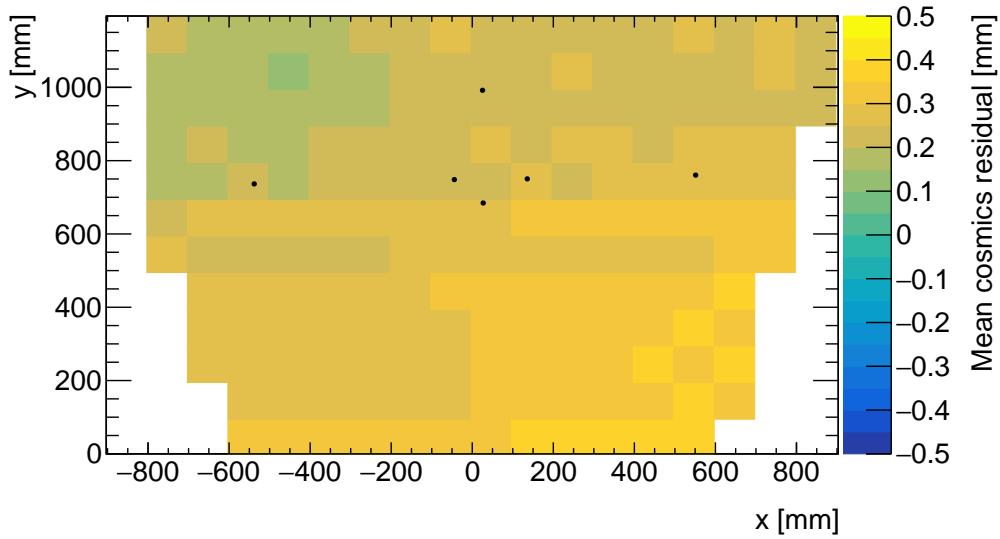
The mean of residuals was extracted for regions across entire quadruplet layers for every tracking combination to get a picture of the how relative misalignments between strip layers change over the layers’ surface. Figure 4.6 contains the mean of residuals on layers 2 and 4 with tracking reference layers 1 and 3 as samples. To understand these plots, realize that the Gaussian mean of the distribution in figure 4.5a is the entry in area bin  $x \in [197, 297]$ ,  $y \in [594.6, 694.6]$  mm in figure 4.6a; likewise for figure 4.5b in figure 4.6b.

Many of the residual means are non-zero, and change smoothly over the layer, indicating that there are relative local offsets stemming from misalignments over larger areas. Given that the residual mean changes with  $x$  in figure 4.6a, there is likely a rotation of layer 2 with respect to layers 1 and 3, combined with an offset of the entire layer. For layer 4 in figure 4.6b, perhaps there is a scaling [8] of the strip pattern with respect to layers 1 and 3. The interpretation of the patterns in the residual means depends on the choice of alignment model.

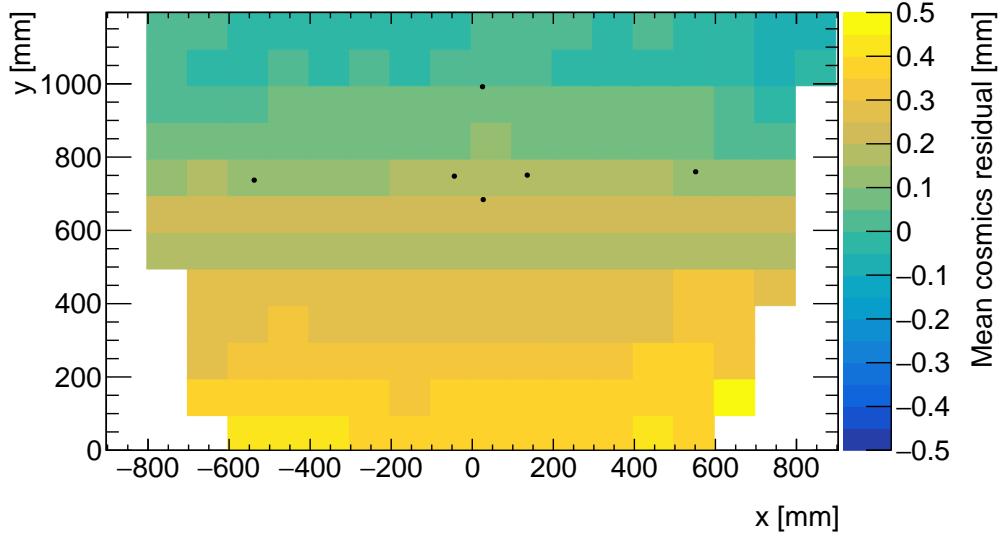
## 4.7 Systematic uncertainty

The statistical uncertainty on the local residual means was typically around 10 - 20  $\mu\text{m}$ , and appendix B shows that the analysis was not statistically limited by the number of triggers collected for each quadruplet. The systematic uncertainties were more significant.

Systematic uncertainties were assigned per tracking combination as the RMS of the distribution of the difference in local residual means each calculated with a different analysis choice.



(a) Mean of residuals for tracks on layer 2, reference layers 1 and 3.



(b) Mean of residuals for tracks on layer 4, reference layers 1 and 3.

Figure 4.6: Mean of residuals in each 100 mm by 100 mm bin over the area of the quad layer for QL2.P.11. The black points represent x-ray survey positions, discussed in chapter 6.

For example, the RMS associated with fitting the local residual distributions with a Gaussian or double Gaussian is 25  $\mu\text{m}$  for the geometrically least favourable tracking combinations. The distribution is shown in appendix C.1. For geometrically similar tracking combinations (like: tracks on layer 1 built from hits on layers 3 and 4, and tracks on layer 4 built from hits on layers 1 and 2), the systematic uncertainty was assigned as the average RMS for both.

Other choices were whether to use data collected at 2.9 kV or 3.1 kV (both are used at McGill); what cluster fitting algorithm to use; and whether or not to apply a differential non-linearity (DNL) correction to the cluster y-positions. A systematic uncertainty was assigned using the method above to account for the effect of each choice. The reasons for each choice are listed below.

Data taken at 3.1 kV was used over 2.9 kV because the strip and wire tracking efficiency increases with higher voltage [36] (appendix C.2).

The `Minuit2` package [42] was used to fit clusters over Guo's method [43] because it provided automatic statistical uncertainty estimates and is the standard fit algorithm choice (appendix C.3).

The DNL correction was not applied because its effect on the residual means was negligible (appendix C.4).

A summary of the systematic uncertainties assigned for each tracking combination is given in table 4.1.

The uncertainty in each mean of residuals was assigned as the sum in quadrature of the statistical uncertainty in the mean and the appropriate systematic uncertainty for the tracking combination.

## 4.8 Discussion

Given that the uncertainty in the mean cosmics residuals is lesser than or near to the order of the required position resolution of the sTGCs (100  $\mu\text{m}$  [5]) the cosmic residual means are relevant input for alignment studies.

Relative local offsets cannot be used to position strips in the absolute ATLAS coordinate system because there was no external reference to measure positions on all layers with respect to. The lack of external reference means that there is not enough information to unfold relative local offsets into absolute local offsets (with respect to the nominal quadruplet geometry). As an example, assuming that the residual on layer 2 in figure 4.4 is representative of the relative local offset, the residual on layer 2 could be caused by the strips on layer 2

Tracking geometry	Residual distribution fit function <a href="#">(C.1)</a>	Cosmics data collection voltage <a href="#">(C.2)</a>	Cluster fit algorithm <a href="#">(C.3)</a>	Apply DNL correction or not <a href="#">(C.4)</a>	Total
Layer 3, fixed layers 1, 2 - like	0.01	0.04	0.02	0.01	<b>0.05</b>
Layer 4, fixed layers 1, 2 - like	0.03	0.01	0.03	0.01	<b>0.10</b>
Layer 2, fixed layers 1, 3 - like	0.01	0.02	0.01	0.000	<b>0.03</b>
Layer 4, fixed layers 1, 3 - like	0.01	0.04	0.01	0.01	<b>0.04</b>
Layer 2, fixed layers 1, 4 - like	0.01	0.04	0.01	0.01	<b>0.04</b>

Table 4.1: Systematic uncertainty assigned for each analysis option, detailed in appendix [C](#).

being misaligned from nominal, but it could also be caused by strips on layers 1 and 4 being offset from nominal while the strips on layer 2 are in their nominal positions! Any number of combinations of local offsets on layers 1, 2 and 4 could produce the residual on layer 2. Absolute local offsets must be calculated another way.

The relative local offsets as calculated from the mean of residual distributions do provide a complete picture of the relative alignment between detectors planes. In fact, cosmic muon testing is the only characterization technique where the entire surface of quadruplet layers can be probed since muons rain down evenly; the CMM [8] and x-ray methods [9] depend on measurements at reference points and test beams only have a limited beam spot [6]. By looking at 2D-histograms of residual means like figure 4.6 for all tracking combinations, it is easy to identify quadruplets that suffer large relative misalignment since many residual means will differ significantly from zero. Moreover, the pattern in the relative local offsets can be used to motivate a physical interpretation of misalignments. Potential application of information is to use it as a reference, cross check, or input in other alignment studies.

# Chapter 5

## Using x-rays to measure relative strip position offsets

Other work on characterizing relative alignments between quadruplet layers has been completed [44] or is ongoing, (Can I cite John's thesis-in-progress?) but what is required are the absolute strip positions with respect to their nominal position in the ATLAS analysis coordinate system to be input into Athena [45]. Somehow, alignment parameters must be derived to create a model of absolute strip positions - which is not possible with the cosmics dataset. Absolute local offset measurements were done with the x-ray method. The reader is referred to the paper describing the x-ray method [9], although the experimental setup has changed slightly since it was written. The experimental setup described here is the most current and was used to collect the data used in this thesis.

### 5.1 Experimental setup

The x-ray tests were performed after the quadruplets arrived at CERN, were assembled into wedges, and alignment platforms installed. Essentially, an x-ray gun was attached to one of the alignment platforms glued to the surface of the wedge and the beam profile recorded by the strips.

The wedges were installed on carts that could rotate their surface to a horizontal position. A mounting platform was installed on top of the alignment platform using a three-ball mount. The x-ray gun used was an [Amptek Mini-X tube](#). The gun was placed in a brass holder with built-in 2 mm collimator and 280  $\mu\text{m}$  copper filter that was mounted on one of five positions on the mounting platform, as shown in figure 5.1. Gun positions were chosen to avoid wire

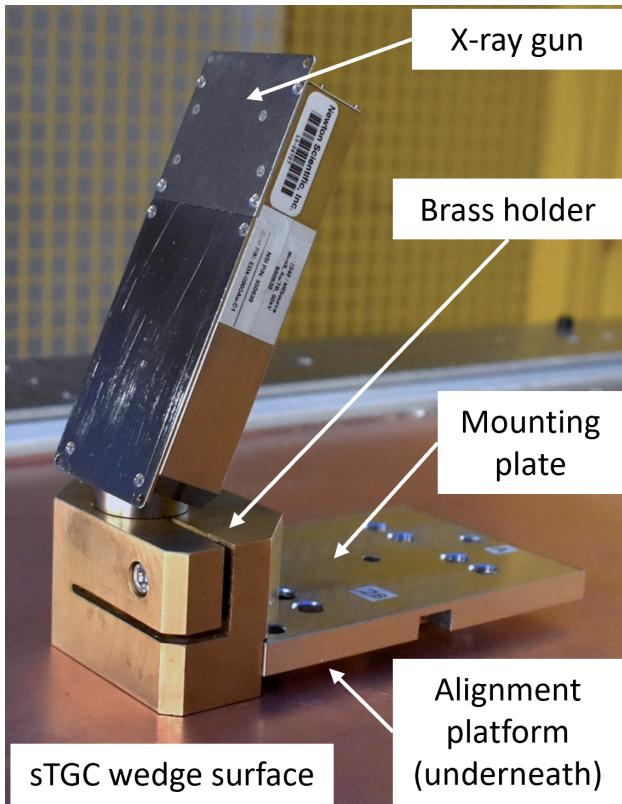


Figure 5.1: The x-ray gun mounted to the alignment platform on the surface of the wedge.  
Adapted from [9].

support structures in the sTGCs that reduce hit efficiency [36] and boundaries between sets of strips readout by two different ASICs that could each have different thresholds.

As with cosmics data collection, each sTGC also needed gas and high voltage to operate. Each layer was operated at 2.925 kV with high voltage from a NIM crate. The chambers were flushed with CO<sub>2</sub> before and during data collection.

The gun produced x-rays with energies under 40 keV with peaks in the 7-15 keV range. The x-rays mostly interacted with the wedge's copper electrodes and gold-plated tungsten wires via the photo effect. The resulting photoelectrons caused ionization avalanches that were picked up by the strips.

## 5.2 Data acquisition

A different version of the same front end electronics, but the same ASIC, as used in cosmics testing were used for the x-ray testing to amplify the data and measure the peak signal amplitude. Data was collected for two minutes per gun position with random triggers. A trigger recorded all signals above threshold. Pad and wire data was not recorded.

## 5.3 Data preparation

Like with cosmics analysis, a default pedestal is subtracted from the signal peak amplitude on each electrode.

Clusters are defined as groups of contiguous strip hits collected within 75 ns. The peak signal amplitude of each electrode in a cluster is fit with a Gaussian, and the mean of the Gaussian is taken as the cluster position. Cluster positions are corrected for DNL (see definition in appendix C.4). Only clusters composed of hits on 3-5 strips were used in the x-ray analysis. Clusters with signal on more than 5 strips were cut because they were most likely caused by photoelectrons ejected with enough energy to be  $\delta$ -rays.

The x-ray analysis must diverge entirely from the cosmics analysis here because the x-rays do not leave tracks. The signals picked up by the strips are from photoelectrons generated on the metals of the sTGCs that only travel through one gas volume and are ejected at all angles. Instead of creating tracks, the cluster position distribution on each layer is used to define the beam profile. A typical beam profile is shown in figure 5.2.

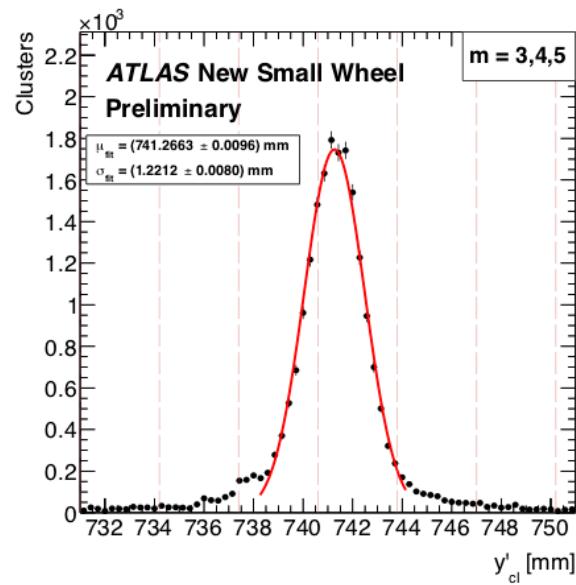


Figure 5.2: Distribution of x-ray cluster mean positions after the analysis cuts and corrections. The strip cluster multiplicity,  $m$ , was limited to 3, 4 and 5. The red line is a Gaussian fit of the distribution and the pink dashed lines denote the edges of the strips [9].

## 5.4 Measuring local offsets

The mean of the cluster position distribution was taken as the x-ray beam profile center. The expected center is calculated assuming a wedge with nominal geometry given the gun position, corrected for: the geometry of the brass holder, the positioning and angle of the alignment platforms and the beam angle. The difference between the expected and reconstructed beam profile center is a measure of the local offset. Applying the logic of equation 4.1 to the beam profile, the fitted mean acts as the recorded position,  $y$ , the expected center is  $y_{nom}$  and the local offset is  $d_{local}$  as before. The x-ray local offsets give the absolute local position of the strip pattern with respect to the source plates. Since the position of the source plates will be monitored by the alignment system in ATLAS [5], the local position of the strip pattern can be known in the ATLAS coordinate system for every position where x-ray data was taken.

The main advantage of the x-ray dataset over the cosmics dataset is that absolute local offsets are measurable thanks to the reference frame provided by the source plates. However, the systematic uncertainty on the x-ray offsets is large: 120  $\mu\text{m}$  was accepted by the collaboration. In addition, local offset measurements were limited to the positions of the alignment platforms; only 10 - 20 positions were surveyed for each wedge. Therefore, validating the x-ray measurements and seeing how they can be improved is important because of the uncertainty in the local offsets and incompleteness of the dataset. Since the local offsets for x-rays and cosmics data are measured in different coordinate systems, they cannot be compared directly. Bringing the cosmics relative local offsets into an absolute coordinate system is impossible; however, the x-ray local offsets can be brought into a relative coordinate system.

## 5.5 Measuring relative local offsets

The measured x-ray beam profile centers were systematically affected by local offsets in the same way as the mean cosmics residuals, as modeled by equation 4.1. Therefore, if a 2-layer track is abstracted from the beam profile centers on each layer and the residual calculated on a third layer, that residual should match the local mean cosmics residual.

The track is “abstracted” because a beam profile center is actually the Gaussian mean of all selected mean cluster positions recorded during the x-ray data taking period. Abstracting a track was necessary because the x-rays cause signal in the chamber via the photoeffect so there were not individual “x-ray tracks” to record. In fact the x-ray data could be collected separately for each layer. Nonetheless, since the effect of local offsets on the beam profile

centers was the same as their effect on the recorded cosmics cluster positions the difference in algorithm between x-ray and cosmics analysis was allowed.

For each x-ray survey position, the x-ray residual was calculated for all possible tracking combinations (which required an x-ray beam profile on at least three layers). The position of the x-ray residuals are shown as black dots over figure 4.6a and 4.6b. Note that the mean of cosmics residuals around the x-ray points were calculated in bins exactly centered on the nominal x-ray gun position, unlike in figure 4.6.I think I want to put copies of the figures I reference here, but with the x-ray residual values labeled. Thoughts Brigitte? It feels like this section needs a "result" plot or table

The uncertainty on the x-ray residuals was the error propagated through the tracking, taking an uncertainty of 120  $\mu\text{m}$  on each beam profile center. The uncertainty on the x-ray residuals ranged from 0.1 mm to 0.4 mm for the most to least geometrically favourable tracking combination - significantly larger than the uncertainty on the mean cosmics residuals.

# Chapter 6

## Comparing cosmic muon and x-ray relative strip position offsets

The goal was to validate the local offsets extracted from the x-ray data with cosmics data. The complication was that the x-ray dataset provided absolute local offsets while the cosmics dataset provided relative local offsets, which could not be compared directly. The solution was to analyze the x-ray data in the same relative coordinate system as the cosmics data. The results of the comparison are presented here.

### 6.1 Assessing correlation

Scatter plots of the x-ray and mean cosmics residuals for two sample quadruplets in figures 6.1 and 6.2 reveal the degree of correlation between the datasets.

First, the fitted slope and offset in figure 6.1 show that the two QL2.P.11 datasets are correlated. However, the magnitude of the uncertainties in the x-ray residuals is large. The large uncertainty set a limit on the sensitivity of the analysis, for if the absolute value of the x-ray residuals of a quadruplet were smaller than the x-ray residual uncertainties, no conclusion about the correlation could be drawn, like for QL2.P.8 (figure 6.2).

Several quadruplets were tested for each quadruplet construction geometry built in Canada. Each quadruplet fell into one of the two categories: residuals large enough to see a correlation, or residuals too small to see a correlation. Since the x-ray and mean cosmics residuals were measures of the relative local offsets between the layer of a quadruplet and the two reference layers, quadruplets with the most relative misalignment had the largest range of residuals.

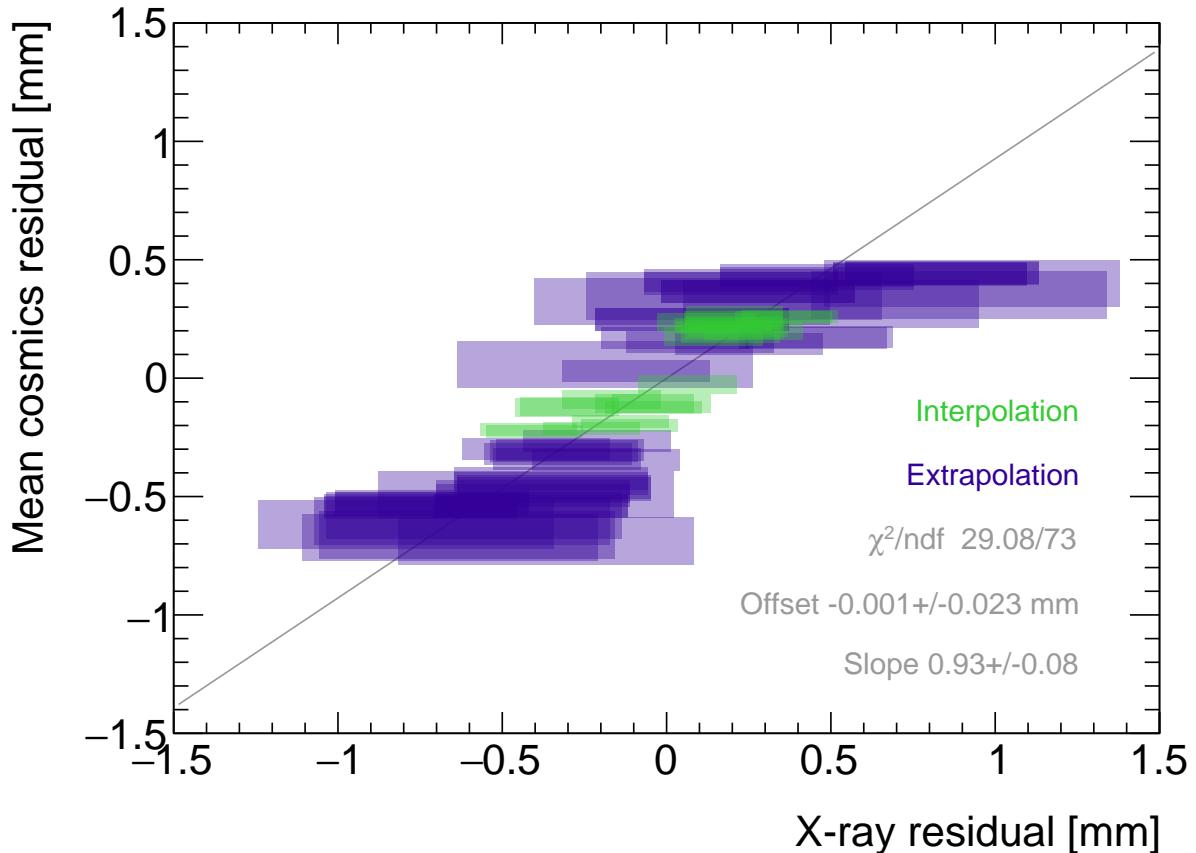


Figure 6.1: Correlation plot between x-ray and mean cosmics residuals for all tracking combinations for QL2.P.11. Each rectangle is centered on an x-ray and mean cosmics residual pair calculated around a certain gun position and for a certain tracking combination. The width of the rectangles in  $x$  and  $y$  are the uncertainty in the x-ray and mean cosmics residual respectively. A printer-friendly version of this plot is available in appendix D.

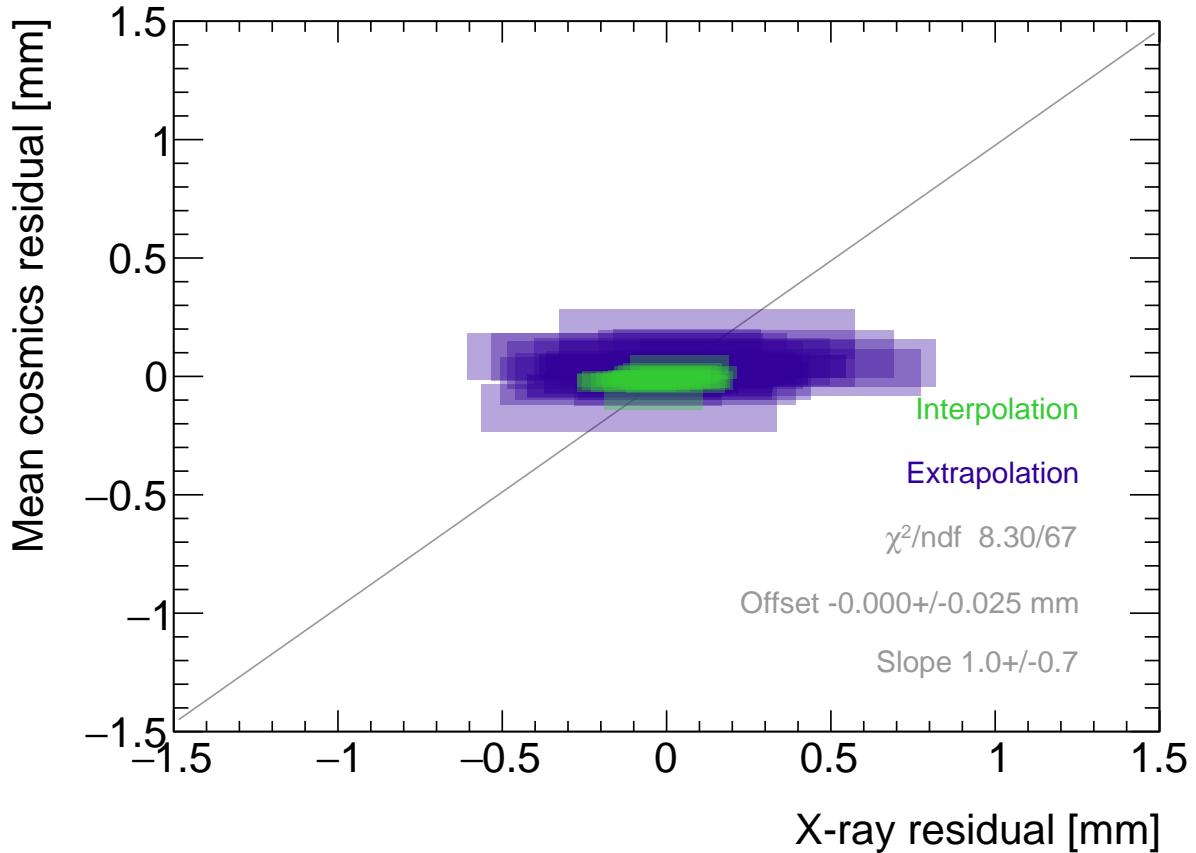


Figure 6.2: Correlation plot between x-ray and cosmics residuals for all tracking combinations for QL2.P.8. Each rectangle is centered on an x-ray and mean cosmics residual pair. The width of the rectangles in  $x$  and  $y$  are the uncertainty in the x-ray and mean cosmics residual respectively. A printer-friendly version of this plot is available in appendix D.

So, the correlation plots were an easy visual way to identify quadruplets with large relative misalignments.

Residuals calculated with all tracking combinations are included in figures 6.1 and 6.2. There are three patterns in the residuals on the scatter plot explained by geometry. First, for both datasets the uncertainty in the extrapolated track residuals were larger than the interpolated track residuals because of the extrapolation lever arm. For the x-ray residuals, the effect of the lever arm on the uncertainty was direct since the residual was calculated from a single abstracted track; for the cosmics residuals it was the widening of the residual distribution on the layer of interest due to the extrapolation lever arm that increased the statistical uncertainty in the fitted mean of residuals and how precisely the effect of systematic choices could be measured. Second, residuals calculated through extrapolation tend to be larger because the extrapolation lever arm can produce more extreme values. Third, the pattern of points in figure 6.1 is slightly mirrored. This is expected since the residuals calculated for a given set of three layers should be geometrically correlated.

The correlation of the cosmics residuals with the x-ray residuals alone does not validate the method; all the studies in described in appendices B and C demonstrate its robustness. The analysis could be validated externally by comparing the mean cosmics residuals to the relative alignment parameters calculated using `tgc_analysis/MatrixMethod` [46] and JOHN FLORES CHI2 METHOD.

## 6.2 Discussion

The most important limit on measuring the degree of correlation between the x-ray and cosmics residuals was the uncertainty on the x-ray residuals, which stemmed from the systematic uncertainty in the x-ray beam profile centers [9]. The method was limited primarily by the sTGCs' poor x-ray position resolution, since x-rays do not create real tracks anyways.

The analysis of certain quadruplets was also limited by the availability of data. Sometimes, less than three layers were surveyed for a given x-ray gun position so no residuals could be calculated. Too few x-ray residuals prevented the analysis from detecting a significant correlation if it would even be measurable. Often, the analysis of smaller quadruplets (placed innermost on the wheel) suffered as a result because they had fewer alignment platforms, and hence gun positions, on their surfaces. In addition, the analysis was limited to certain quadruplets. The wedges constructed the earliest (typically small wedges) were surveyed when the method was still being designed and so have limited x-ray residuals calculated from beam profiles of lower quality. Also, not all cosmic muon test sites had enough front

end electronics to collect data on three layers simultaneously, which is the minimum required to be able to calculate residuals.

Nonetheless, the comparison of x-ray and cosmics residuals was really to confirm the x-ray method's ability to measure local offsets with an independent dataset. The x-ray absolute local offsets allow the calculation of relative local offsets that can be correlated to the cosmics relative local offsets. Therefore, the analysis of quadruplets with relative local offsets large enough to detect a correlation validates the x-ray method's ability to measure local offsets. The potential of using relative local offsets calculated from cosmics data to study relative alignment between sTGC layers stands on its own. It is especially compelling considering how much smaller the uncertainties in the relative local offsets are compared to the x-ray data.

# Chapter 7

## Outlook and summary

The cosmic muon dataset was used to independently confirm the absolute local offsets measured by the x-ray method. The x-ray offsets are being used to complete the sTGC alignment scheme of the NSWs: the NSW alignment system monitors the position of alignment platforms on the surface of sTGC wedges, and the x-ray measurements provide the offsets of the strip pattern with respect to each alignment platform. The continuation of this analysis is detailed next (section 7.1) before considering the larger context (section 7.2).

### 7.1 Outlook

Next all quadruplets with suitable cosmics and x-ray data should be surveyed to flag anomalous quadruplets (as a first step). If a quadruplet’s correlation plot like figure 6.1 or 6.2 reveals an unexpected correlation or has a large scatter, it would indicate an issue with either the cosmics or x-ray data collection to be investigated further. The uncertainty in each set of tracking points would inform the interpretation of the anomaly. Then, the quality of the correlation should be evaluated over all quadruplets instead of individually.

For now, the correlation for the individual quadruplets tested support the use of the x-ray data to build a global alignment model [9]. Work on creating an alignment model is ongoing with the development of `stgc_as_built_fit` [47]. Currently, the algorithm compares the y-position of a local group of strips at each x-ray gun position as measured by the x-ray and CMM methods in a fit to extract a global slope ( $m$ ) and offset ( $b$ ) per layer,  $l$ , where the  $\chi^2$  is given by equation 7.2.

$$dy_{cmm,corr} = y_{cmm} + b_l + m_l x - y_{nom} \quad (7.1)$$

$$\chi^2 = \frac{[dy_{cmm,corr} - dy_{xray}]^2}{\delta dy_{xray}^2 + \delta dy_{cmm,corr}^2} \quad (7.2)$$

Here,  $dy$  refers to the corrected CMM and x-ray local offsets, and  $\delta dy$  refers to their respective uncertainties. The CMM measurements were taken before the cathode boards were assembled into quadruplets, so alignment parameters for the given layer were extracted from the  $\chi^2$  fit by stepping the corrected CMM y-position towards the x-ray y-position by adjusting the alignment parameters. The plan is that the alignment parameters will be provided to [Athena](#) [45] to precisely reconstruct muon tracks from the NSWs' sTGCs. The large uncertainty on the x-ray local offsets ( $120\text{ }\mu\text{m}$ ) and the sparseness of the measurements means that including input from other characterization datasets could reduce the uncertainty alignment model parameters.

The uncertainty in the mean cosmics residuals, the measure of relative local offsets, was smaller than the desired position resolution of the sTGCs, so they provide relevant information about strip positions. Moreover, they can be calculated over the entire area of the quadruplet instead of at specified positions. It would be great to use the cosmics residuals as input to calculate and reduce the uncertainty on the alignment parameters. Since mean cosmics residuals can only provide relative alignment information, one idea would be to use them to constrain the fit of the alignment parameters. In this case, the alignment parameters would need to be fitted on all layers at once, and the shifting y-positions on each layer forced to create an abstracted track residual equal to the local mean cosmics residual (within uncertainty) for each x-ray point. Or, instead of constraining the fit, it could be penalized if the resulting parameters do not result in abstracted track residuals equal to the mean cosmics residuals within uncertainty. Some work on using the three datasets at once in a fit has been started.

## 7.2 Summary

The LHC [2] will be at the energy frontier of particle physics for many years to come, making it a unique tool with which to study particle physics. With the HL-LHC [1], high statistics on rare particle physics processes will enable more precise measurements of parameters of the Standard Model and increase the sensitivity to signatures of physics beyond the Standard Model [3]. To capitalize on the increased collision rate, the NSWs of the ATLAS experiment must be replaced to keep good triggering and tracking performance [5].

Small-strip thin gap chambers are gas ionization chambers optimized for a high rate environment [5]. Using the pad electrodes to define a region of interest makes it possible to get

track segments of  $\sim 1$  mrad angular resolution quickly, which will be used as input to check if a collision originated from the interaction point and should be triggered on or not [5, 7]. sTGCs are also able to provide better than 100  $\mu\text{m}$  position resolution on each detector plane to fulfill precision offline tracking requirements [6].

Ultimately, the positions of the sTGC strip electrodes need to be known in ATLAS to within  $\sim 100 \mu\text{m}$  so that they can deliver the required position resolution. The ATLAS alignment system will position alignment platforms on the surface of the sTGC wedge, and an alignment model will be used to position the strips with respect to the alignment platforms [5]. Input to the alignment model comes from the datasets used to characterize the quadruplets. The x-ray method [9] is used to measure offsets of strips from their nominal position to achieve this goal. The alignment model could be built on x-ray data alone, but the sparseness of and large uncertainty on the local offsets mean that the alignment model could benefit from more input. Comparing the x-ray offsets to the CMM data [8] allows the effect of inter-layer misalignments to be isolated and increases the input to the alignment model.

The cosmics dataset was used to confirm the local offsets measured with the x-ray gun. It provides relative local offsets between sTGC strip layers. The 2D visualizations of relative local offsets allow personnel to quickly identify areas of misaligned strips and make hypotheses of the physical origin of those misalignments. The correlation seen between the x-ray and cosmics relative local offsets in quadruplets with large relative misalignments both confirms the validity of the x-ray local offsets and again is a quick way to identify quadruplets with anomalous misalignments. Moreover, the mean of track residuals in an area is a robust estimation of the relative local offset, as shown by the estimation of systematic uncertainties; the relative local offsets for all 2-fixed layer reference frames do not change by more than 100  $\mu\text{m}$  given variation in data collection conditions and analysis algorithms. The cosmics relative local offsets are therefore relevant input for alignment studies and could improve the alignment model that will position each strip.

Achieving the required position resolution on each layer of the NSWs in the particle track bending plane achieves the design momentum resolution for muons ejected towards the end-caps of ATLAS. Muons are important signatures of electroweak and Higgs sector events of interest for the ATLAS Collaboration’s future physics goals [5]. Being the second of two tracking technologies on the NSWs, an effective alignment model of sTGC quadruplets layers is a necessary part of making the NSWs redundant for 10 or more years of recording collisions in the High Luminosity era of the LHC.

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# APPENDICES

# Appendix A

## Uncertainty in cluster positions

### A.1 Cluster definition

A cluster is a series of contiguous strip channels on a layer with non-zero amplitude, all part of the same trigger and having the same event number [36]. Clusters result from the drift of ionization products generate in the ionization avalanche caused by a muon [29]. The peak-detector-output (PDO) of the signal on each strip of a cluster is fit with a Gaussian. The y-position of a particle as it passed through the layer is mean of the cluster, referred to here as the hit position.

### A.2 Effect of fit algorithm on cluster mean

The clusters were fit with Guo's method [43] and Minuit2 for ROOT [42]. The difference in cluster means between the two algorithms is shown in figure A.1.

The RMS of the distribution in figure A.1 is 57  $\mu\text{m}$ , which is much larger than the statistical uncertainty in the mean for the Minuit2 algorithm, which peaks around 7  $\mu\text{m}$ . An RMS of 60  $\mu\text{m}$  is common for data taken with most quadruplets at 3.1 kV. Therefore, the uncertainty in the y-hit positions is assigned 60  $\mu\text{m}$ .

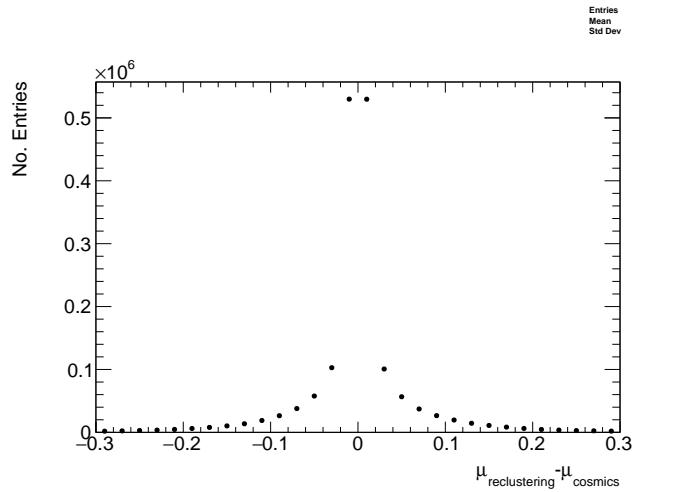


Figure A.1: The difference between cluster means calculated with Guo's method [43] in `tgc_analysis/CosmicsAnalysis` and `Minuit2` for `ROOT` [42] in `strip_position_analysis/ReClustering` for data collected with QL2.P.8 at 3.1 kV.

### A.3 Effect of uncertainty in cluster mean on track residuals

The uncertainty assigned to the hit position affected the uncertainty in the extrapolated/interpolated position of the track, and in the residuals. The bin size of the residual distributions was set to 200  $\mu\text{m}$  because that was the uncertainty in the residuals calculated from the tracks with the least favourable geometry (like tracks built from hits on layers 1 and 2 and extrapolated to layer 4).

## Appendix B

### Study of cosmics for alignment analysis statistical uncertainty

Typically, one million triggers (cosmic muon events and noise) were collected for each Canadian quadruplet at McGill University, resulting in roughly half the number of viable tracks after cuts in `tgc_analysis/CosmicsAnalysis`. For QS3.P.18, 3.5 million triggers were collected. To gauge the sensitivity of the analysis to the available statistics, partitions of this data with each with a different number of triggers were analyzed separately. Ultimately, the quantity of interest was the gaussian mean of the residual distribution in regions of interest, so the peak in the distribution of the statistical uncertainty in the residual means for each area of interest for a specific tracking combination was used to gauge the quality of the analysis. How the peak in the residual mean uncertainty distribution changes with the number of triggers is shown in figure for tracks on layer 1 built from layers 3 and 4 [B.1](#).

The uncertainty is already around  $20\text{ }\mu\text{m}$  at 1 million triggers, suitable for distinguishing differences in offsets of order  $50\text{ }\mu\text{m}$  as required. Although increased statistics could decrease the statistical uncertainty, it is not required for the goals of this analysis. Moreoever, the systematic uncertainty is around  $50\text{ }\mu\text{m}$  and the systematic uncertainty on the x-ray residuals is  $150\text{ }\mu\text{m}$  so the statistical uncertainty of  $20\text{ }\mu\text{m}$  is nearly negligible.

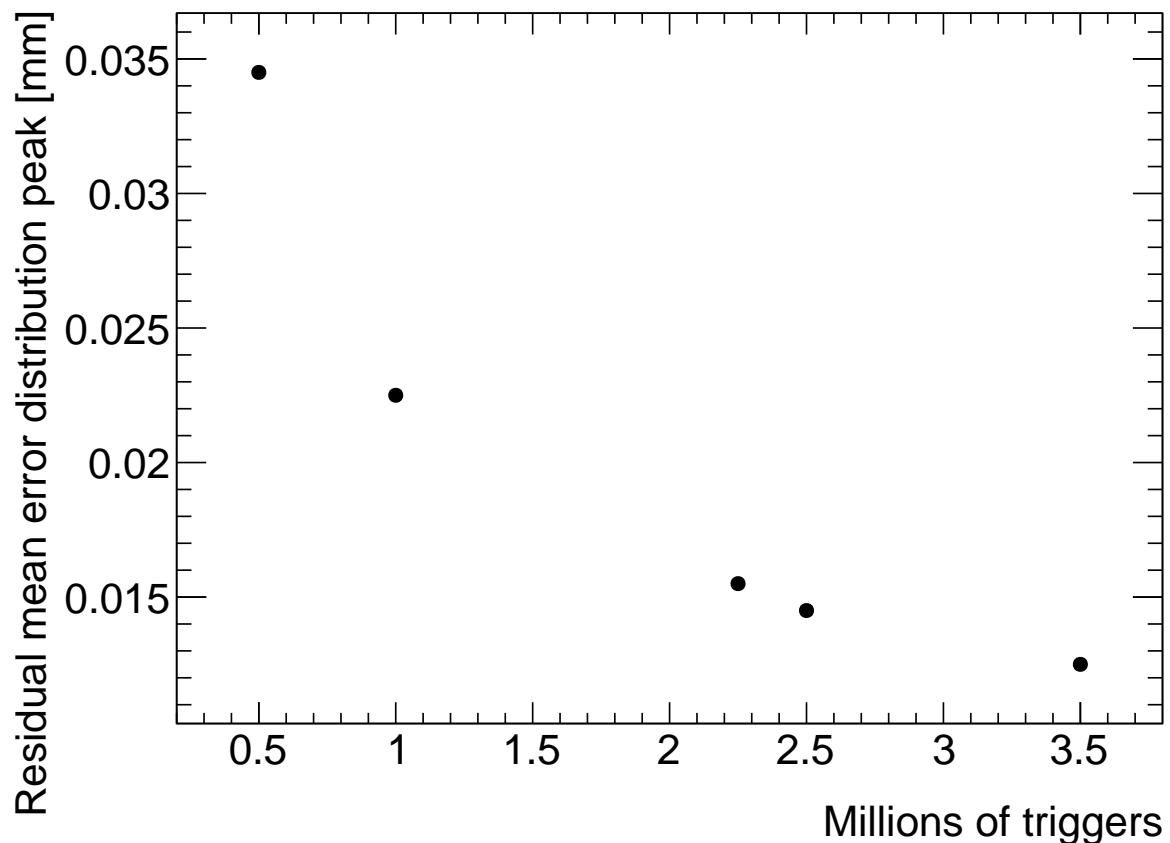


Figure B.1: How the peak of the distributions of uncertainties in the residual means in regions of interest for tracks on layer 1 built from layers 3 and 4 changed with the number of triggers used in the analysis. The distribution falls off as  $\frac{1}{\sqrt{N}}$  as expected.

# Appendix C

## Study of systematic uncertainties when using cosmics data for alignment studies

### C.1 Residual distribution fit function

The distribution of residuals should be modelled by a double gaussian fit[36]:

$$G(r) = A_s \exp\left[\frac{-(r - \mu)^2}{2\sigma_s^2}\right] + A_b \exp\left[\frac{-(r - \mu)^2}{2\sigma_b^2}\right] \quad (\text{C.1})$$

where  $r$  is the residual,  $A$  is the gaussian amplitude,  $\mu$  is the gaussian mean,  $\sigma$  is the gaussian sigma, and the subscripts  $s$  and  $b$  stand for signal and background respectively. One gaussian captures the real (signal) tracks and the other captures the tracks built from noise (background). The gaussian with the smaller width is identified as the signal.

A single gaussian fit failed less often than a double gaussian fit. The gaussian fits were performed by initially estimating the amplitude to be 100 tracks, the gaussian mean to be the histogram mean, and gaussian  $\sigma$  to be the RMS. The fit range was restricted to  $\pm 1$  RMS from the histogram mean. The modification helped the gaussian fit capture the signal peak. An example residual distribution is shown in figure C.1.

For all residual distributions in 100 mm by 100 mm bins on layer 4 built from hits on layers 1 and 2, the difference in gaussian and double gaussian means and  $\sigma$ 's is shown in figure C.2. Since the RMS of the residual mean differences distribution is less than 50  $\mu\text{m}$  the gaussian

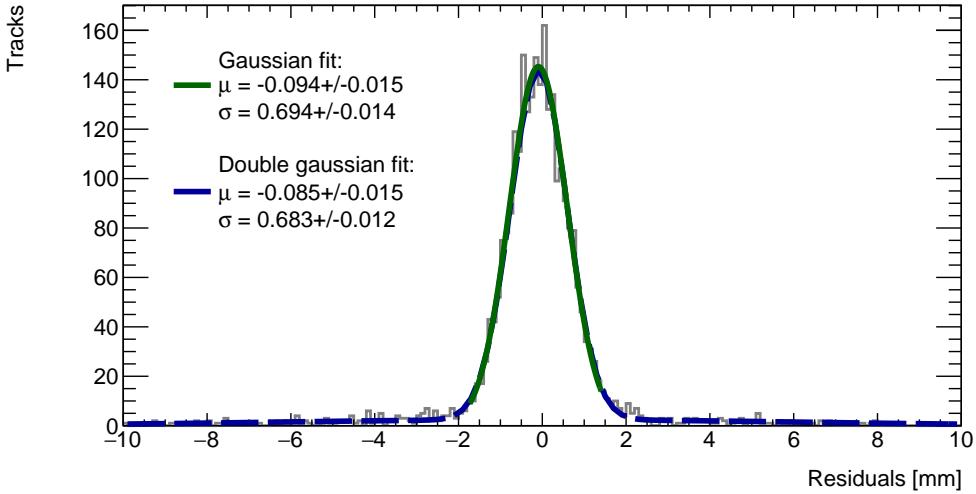


Figure C.1: Residual distribution for tracks on layer 4 built from hits on layers 1 and 2 for  $x \in [-3.00, 97.00]$ ,  $y \in [394.60, 494.60]$  mm for QL2.P.8 fit with a double gaussian and a single gaussian in a range of  $\pm 1$  RMS from the histogram mean.

fit gave the same result within the required precision. Moreover, this is for the tracking combination with the worst extrapolation lever arm and the widest distribution of mean differences; the interpolation combinations have narrower distributions.

The gaussian  $\sigma$  should be larger than the double gaussian  $\sigma$  because the gaussian distribution includes the effect of the noise tracks with large residuals, while the double gaussian models signal and background residuals separately. For this analysis, only the residual mean was important, so the systematic overestimate of the signal  $\sigma$  in the gaussian fit shown on the right of figure C.2 was allowed.

## C.2 Cosmic muon data collection voltage

Cosmic muon data was collected at 2.9 kV and 3.1 kV because although 2.9 kV is closer to the operating conditions the chambers will be subject to in ATLAS, the extra gain provided by operating at 3.1 kV increased the signal to noise ratio for pad signals. Also, the tracking efficiency was higher with data collected at 3.1 kV. The difference in gain affected the relative population of clusters of different sizes, which in turn affected the uncertainty in the strip hit positions on each layer, the uncertainty in the track positions and the residual distributions. The residual distributions for 3.1 kV data are narrower, as shown in figure C.3.

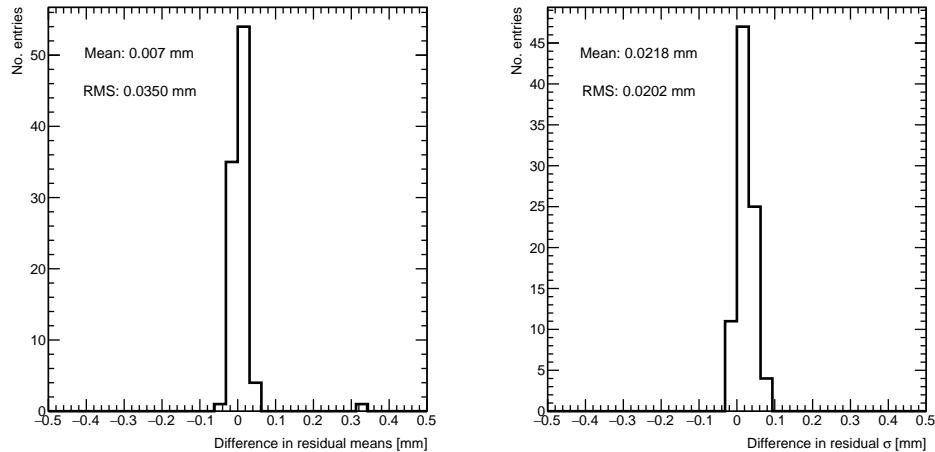


Figure C.2: Difference in residual distribution means and  $\sigma$ 's for a gaussian and double gaussian fit, for all residual distributions in 100 mm by 100 mm bins on layer 4 built from hits on layers 1 and 2 for QL2.P.8.

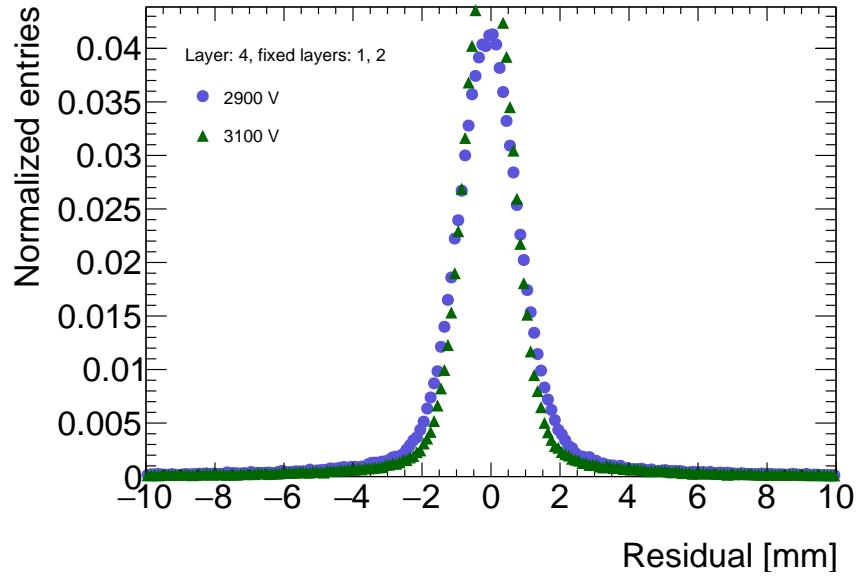


Figure C.3: Residual distribution for tracks on layer 4 built from hits on layers 1 and 2 for QL2.P.8 for data collected at 2.9 kV and 3.1 kV.

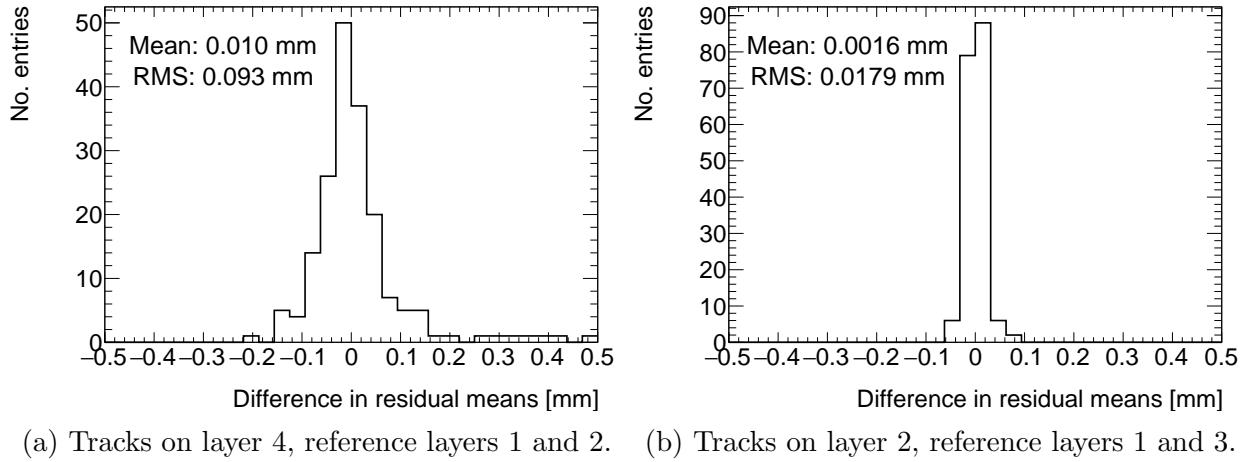


Figure C.4: Difference in residual means for data collected with QL2.P.8 at 2.9 kV and 3.1 kV respectively in 100 mm by 100 mm bins for (a) tracks on layer 4 built from hits on layers 1 and 2 and (b) tracks on layer 2 built from hits on layers 1 and 3.

Neither dataset is better for calculating the mean of residuals in a given area, so a systematic uncertainty can be assigned based on the difference in residual means calculated for 2.9 kV and 3.1 kV data; namely, the systematic uncertainty was approximated as the RMS of the residual mean difference distribution. Data taken with QL2.P.8 was used to estimate the RMS, as in figure C.4a.

Tracks built from hits on layers 1 and 2 and extrapolated to layer 4 have the worst lever arm and hence the most uncertainty. The width of the distribution for geometrically favourable tracks are much narrower. The narrowest width of the residual mean difference distribution is for tracks on layer 2 built from hits on layers 1 and 3 (see figure C.4b).

Therefore, for each tracking combination, a systematic uncertainty equal to the RMS of the residual mean difference distribution was assigned.

### C.3 Cluster fit algorithm

To ensure that changing the cluster fitting algorithm like in appendix A would not change the calculated mean of residuals in each region of interest significantly, the residual means were compared in both cases. The distribution of the difference in residual means is plotted in figure C.5 for the tracking combination with the worst extrapolation lever arm.

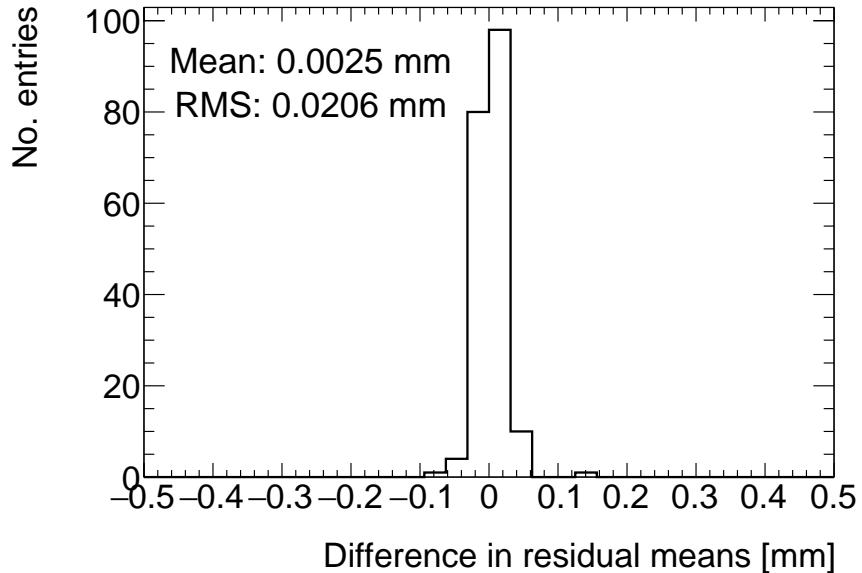


Figure C.5: Difference in residual means when the cluster fit algorithm is `Minuit2` [42] versus Guo’s method [43] for tracks on layer 4 built from hits on layers 1 and 2 for QL2.P.8.

The other tracking combinations had smaller RMS values. Differences on the order of 50  $\mu\text{m}$  are important, so figure C.5 shows that the clustering algorithm had a small but notable effect. Therefore, the RMS for each tracking combination will be used to add a systematic uncertainty on the residual means.

## C.4 Differential non-linearity

### Definition

In this context, differential non-linearity (DNL) is when the reconstructed cluster mean is biased by the fit of the discretely sampled PDO distribution over the strips. The bias depends on the relative position of the avalanche with respect to the center of the closest strip. For a summary of DNL, refer to page 40 of Lefebvre’s thesis [36] and for an example application, refer to [6].

### Application and effect of DNL

The cluster mean was corrected for DNL using the equation:

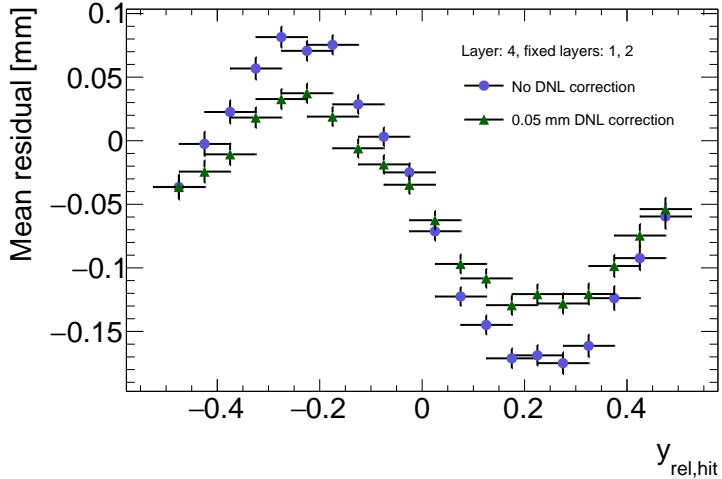


Figure C.6: Effect applying a 50  $\mu\text{m}$  DNL correction to the cluster means on the residual vs  $y_{rel}$  distribution for tracks built from layers 1 and 2 and extrapolated to layer 4 for QL2.P.8.

$$y' = y + a \sin(2\pi y_{rel}) \quad (\text{C.2})$$

where  $y$  is the cluster mean,  $y_{rel}$  is the relative position of the cluster mean with respect to the strip's center,  $a$  is the amplitude of the correction, and  $y'$  is the corrected cluster mean. The amplitude can be derived by comparing the reconstructed hit position to the expected hit position, as done in Abusleme, 2016 [6]. With cosmic muons, there is no reference hit position to compare to, so track residuals were used as a proxy [36]. The hallmark of the DNL effect is the periodic pattern in the residual versus  $y_{rel}$  profile, and the effect of correcting the cluster means using an amplitude of 50  $\mu\text{m}$  is shown in figure C.6. An amplitude of 50  $\mu\text{m}$  was based on Lefebvre's estimate of the DNL amplitudes by layer, quadruplet and cluster size using exclusive cosmic muon tracks in `tgc_analysis/CosmicsAnalysis`. Little variation was seen in the amplitude parameters with respect to the quadruplet tested, the layer and the cluster size so a universal correction was used.

Although the correction is not large enough in this case, the figure shows that the correction does reduce the DNL effect. Slightly better performance is seen in the interpolation tracking combinations where the quality of the residuals is better. DNL corrections for cosmic muon data are difficult because the DNL effect is obscured by the effect of misalignments and noise. Misalignments cause the center of the sine pattern in figure C.6 to be shifted off of zero, since the mean of residuals is shifted.

In figure C.7, it is apparent that the effect of the DNL correction on the mean of the

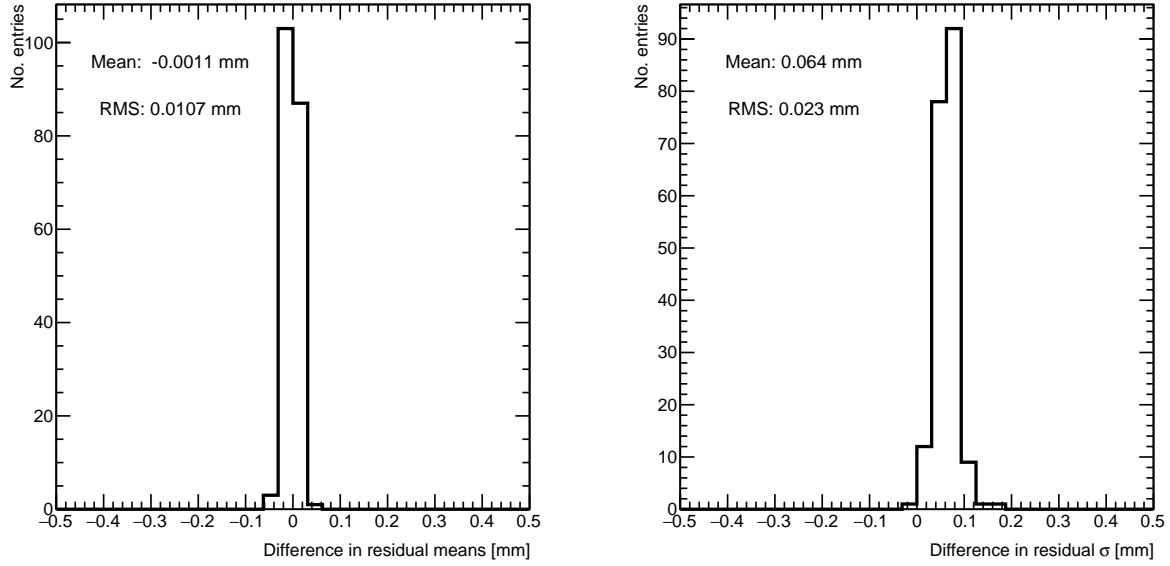


Figure C.7: Difference in residual distribution means and  $\sigma$ 's with and without DNL correction for residuals on layer 4 from reference layers 1 and 2 for QL2.P.8.

residual distribution in 100 mm by 100 mm areas is on the order of micrometers in the worst extrapolation case. Although the  $\sigma$ 's of the residual distributions shrink with the DNL correction, the mean is the parameter of interest. Therefore, for this analysis DNL was not corrected for.

The  $\sigma$ 's of the residual distributions do shrink with the DNL correction but not so much to affect the residual means, which are the important parameter for this analysis. Therefore, since the effect of the DNL correction is negligible, it was not pursued further.

## Appendix D

### Printable plots

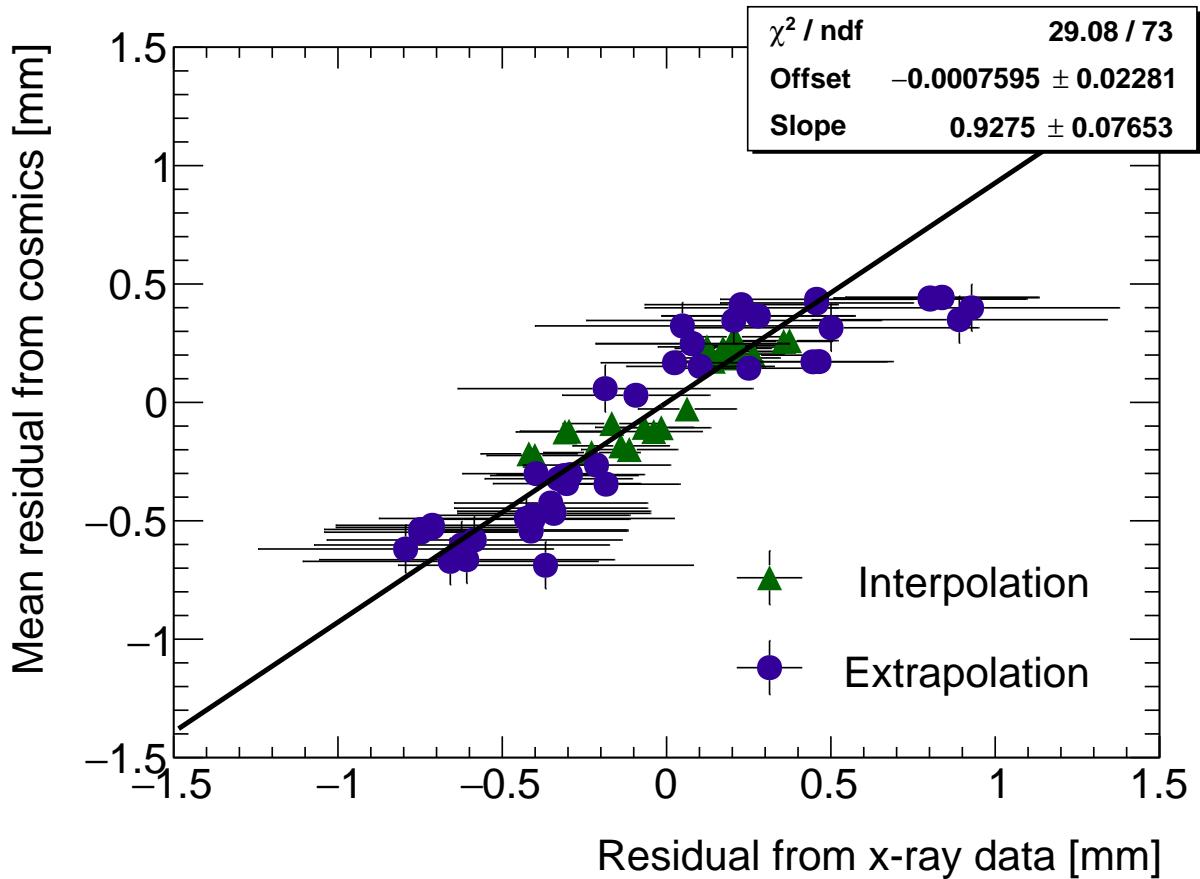


Figure D.1: Correlation plot between x-ray and cosmics residuals for all tracking combinations for QL2.P.11. Each rectangle is centered on an x-ray and mean cosmics residual pair. The width of the rectangles in  $x$  and  $y$  are the uncertainty in the x-ray and mean cosmics residual respectively. This is a printable version of figure 6.1 in section 6.1.

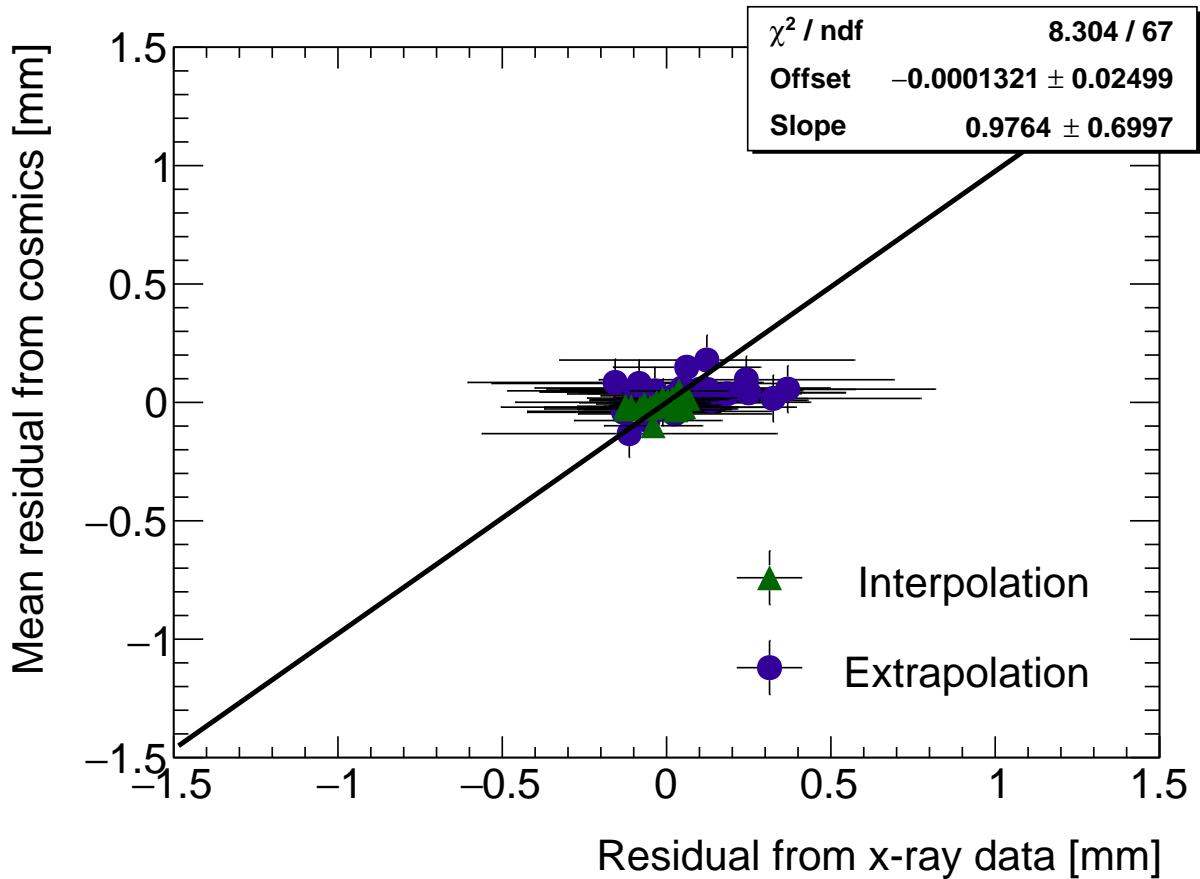


Figure D.2: Correlation plot between x-ray and cosmics residuals for all tracking combinations for QL2.P.8. Each rectangle is centered on an x-ray and mean cosmics residual pair. The width of the rectangles in  $x$  and  $y$  are the uncertainty in the x-ray and mean cosmics residual respectively. This is a printable version of figure 6.2 in section 6.1.