



# Dynamical Timescale of Pre-collapse Evolution Inferred from Chemical Distribution in the Taurus Molecular Cloud-1 (TMC-1) Filament

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## Abstract

We present observations and analyses of the low-mass star-forming region, Taurus Molecular Cloud-1 (TMC-1). CS ( $J = 2-1$ )/N<sub>2</sub>H<sup>+</sup> ( $J = 1-0$ ) and C<sup>17</sup>O ( $J = 2-1$ )/C<sup>18</sup>O ( $J = 2-1$ ) were observed with the Five College Radio Astronomy Observatory and the Seoul Radio Astronomy Observatory, respectively. In addition, *Spitzer* infrared data and 1.2 mm continuum data observed with Max-Planck Millimetre Bolometer are used. We also perform chemical modeling to investigate the relative molecular distributions of the TMC-1 filament. Based on *Spitzer* observations, there is no young stellar object along the TMC-1 filament, while five Class II and one Class I young stellar objects are identified outside the filament. The comparison between column densities calculated from dust continuum and C<sup>17</sup>O 2–1 line emission shows that CO is depleted much more significantly in the ammonia peak than in the cyanopolyyne peak, while the column densities calculated from the dust continuum are similar at the two peaks. N<sub>2</sub>H<sup>+</sup> is not depleted much in either peak. According to our chemical calculation, the differential chemical distribution in the two peaks can be explained by different timescales required to reach the same density, i.e., by different dynamical processes.

**Key words:** astrochemistry – ISM: abundances – ISM: individual objects (TMC-1) – ISM: molecules

## 1. Introduction

Gravitationally bound starless cores, that is, prestellar cores, are cold ( $T_k \sim 10$  K), dense ( $n > 10^4 \text{ cm}^{-3}$ ) condensations of gas with no sign of a central protostellar object, implying that they are the earliest identifiable stage of star formation (e.g., Shu et al. 1987; Lee & Myers 1999; Crapsi et al. 2005; Ward-Thompson et al. 2007). Therefore, prestellar cores are places in which we can study the initial conditions of star formation such as the distribution of density, temperature, and velocity (Benson & Myers 1989; Ward-Thompson et al. 1994). Prestellar cores have a constant density in the inner part and a decreasing density at larger radii (Ward-Thompson et al. 1994, 1999, 2000; André et al. 1996), which can be well described with a Bonnor–Ebert sphere density profile (Evans et al. 2001). Dust grain temperatures decrease toward the center of prestellar cores (Evans et al. 2001; Zucconi et al. 2001; Ward-Thompson et al. 2002; Crapsi et al. 2007; Harju et al. 2008; Nielbock et al. 2012) since the interstellar radiative field (ISRF) is the only heating sources of prestellar cores.

Although there are several model scenarios to explain core formation, they can be divided into two extreme paradigms: slow quasistatic contraction and dynamic turbulent processes. In the former process, a core gradually becomes condensed toward the center through ambipolar diffusion, increasing the ratio of core mass to magnetic flux. As a core switches from magnetically subcritical to a supercritical condition, the core can collapse gravitationally (Shu et al. 1987; Mouschovias 1991; Mouschovias & Ciolek 1999). If compressible turbulence dominates in a prestellar core, the turbulence will decline in a short time (e.g., Mac Low et al. 1998) and collapse will happen. This means that the prestellar core disappears quickly or rapidly collapses. The detailed modeling results of this

process suggests that the cores live for only one or two free-fall times (e.g., Ballesteros-Paredes et al. 2003; Vázquez-Semadeni et al. 2005). The demographic studies indicate lifetimes that decline with increasing density from near 10 free-fall times at  $10^3 \text{ cm}^{-3}$  to a few free-fall times, or about 0.5 Myr, once the mean density rises above  $2 \times 10^4 \text{ cm}^{-3}$ , and one free-fall time once the mean density is above  $10^6 \text{ cm}^{-3}$  (Ward-Thompson et al. 2007; Enoch et al. 2008; André et al. 2014).

Studies show that the chemical compositions of dark-cloud cores can be a useful tracer of their evolutionary status in a given dynamical process because the chemistry is time dependent. The dynamical evolutionary stage can be inferred by the density structure of a core, which can be obtained through the dust continuum emission. However, one of the important properties of dynamical processes is the velocity distribution, which is traced only by line profiles within the core. Molecular line profiles are affected by not only density and temperature profiles but also the abundance structure in the core. Therefore, it is necessary to understand the chemical distribution within cores (e.g., Mizuno et al. 1990; Zhou et al. 1993; Bergin et al. 1995; Choi et al. 1995; Gregersen et al. 1997; Mardones et al. 1997; Gregersen & Evans 2000; Aikawa et al. 2001; Di Francesco et al. 2001; Lee et al. 2003, 2004; Belloche et al. 2002).

Taurus Molecular Cloud-1 (TMC-1) is an interesting source for exploring the chemistry of dark clouds. It is a cold and dense cloud with a narrow filamentary structure ( $5' \times 15'$ ) in the Taurus dark-cloud complex at a distance of  $\sim 140$  pc (Elias 1978; Kenyon et al. 1994; Torres et al. 2012) extending from the southeast to the northwest direction, with an infrared source (IRAS 04381+2540) located outside the northwest part. The TMC-1 filament has been observed in many molecular lines. According to those molecular observations, carbon-chain

molecules such as CCS, C<sub>4</sub>H, and HC<sub>3</sub>N are abundant in the southeast part (cyanopolyne peak) of the TMC-1 filament, and NH<sub>3</sub> and N<sub>2</sub>H<sup>+</sup> are abundant in its northwest part (ammonia peak). This peculiar chemical feature of TMC-1 has been an interesting problem, so there have been many studies attempting to understand the chemical distributions in TMC-1 (e.g., Churchwell et al. 1978; Little et al. 1979; Wootten et al. 1980; Toelle et al. 1981; Guélin et al. 1982; Snell et al. 1982; Schloerb et al. 1983; Olano et al. 1988; Hirahara et al. 1992; Suzuki et al. 1992; Pratap et al. 1997; Hirota et al. 1998, 2003; Markwick et al. 2000, 2001; Saito et al. 2002; Suutarinen et al. 2011).

Many studies have concluded that the cyanopolyne peak is at an earlier evolutionary stage than the ammonia peak, and thus the cyanopolyne peak (southeast part) is less chemically evolved than the northwest part based on the chemical models and molecular line observations such as CCS, HC<sub>5</sub>N, C<sup>34</sup>S, NH<sub>3</sub>, C<sup>34</sup>S, DCO<sup>+</sup>, H<sup>13</sup>CO<sup>+</sup>, and CH. The estimated chemical evolutionary age difference between the two peaks is more than 10<sup>5</sup> years (e.g., Hirahara et al. 1992; Suzuki et al. 1992; Hanawa et al. 1994; Howe et al. 1996; Pratap et al. 1997; Saito et al. 2002; Suutarinen et al. 2011).

Hirahara et al. (1992) showed that the H<sub>2</sub> density of the southeast part was lower than that of the northwest part by a factor of 10 using C<sup>34</sup>S. They suggested that the cyanopolyne peak was in an earlier stage of chemical evolution than the ammonia peak because of the higher density at the ammonia peak. Pratap et al. (1997) observed several molecular species and suggested two possibilities for the chemical difference along the TMC-1 filament. One was a small variation in the C/O ratio from 0.4 to 0.5. Another was that differences in density, which varies by a factor of two along the filament, could explain the difference of chemical distribution in the TMC-1 filament. The latter was consistent with the result by Hirahara et al. (1992) because the higher density could drive chemical processes more rapidly. Recently, Suutarinen et al. (2011) presented a CH abundance gradient in TMC-1, from  $\sim 1.0 \times 10^{-8}$  in the northwestern part to  $\sim 2.0 \times 10^{-8}$  in the southeastern part of TMC-1, and they noted that the CH column density peaks close to the cyanopolyne peak. They suggested that the southeastern part in the vicinity of the cyanopolyne peak had an extensive low-density envelope based on the comparison of the dust continuum maps from SCUBA 850 and 450  $\mu$ m (Nutter et al. 2008) and the modeling of molecular lines (Hirahara et al. 1992; Pratap et al. 1997), implying that the cyanopolyne peak was in the early stage of dynamical evolution and this peak was also chemically less evolved than the ammonia peak because of the lower density.

Markwick et al. (2000, 2001) proposed another explanation for molecular abundance differences along the TMC-1 filament. They investigated the possibility that the gradients were produced by spontaneous explosive desorption of UV-photolyzed ice mantles, triggered by the heating of dust grains and their ices, through grain–grain collisions induced by MHD waves from IRAS 04381+2540 near the ammonia peak. The chemical processes occurred first in the ammonia peak and later in the cyanopolyne peak, which induced the gradient of molecular abundances along the TMC-1 filament. According to their chemical models, the ammonia peak is  $1.5 \times 10^5$  years older than the cyanopolyne peak, assuming the speed of the MHD wave of 2 km s<sup>-1</sup>.

A very different explanation for the chemical distribution along the TMC-1 filament might be possible. According to gas-phase chemical models, carbon-chain species such as CCS form before atomic carbon mainly reacts with atomic oxygen, which drives all carbon-bearing species into CO. When the depletion of CO (and even N<sub>2</sub>H<sup>+</sup>) is very significant, abundances of carbon-chain species increase again, which is called a late-time secondary abundance peak of carbon-chain species (Ruffle et al. 1997, 1999; Li et al. 2002; Lee et al. 2003). In this case, the cyanopolyne peak might be chemically older than the ammonia peak, unlike the conclusion from the rest of the studies. Therefore, we need to calculate accurate depletion factors of N<sub>2</sub>H<sup>+</sup> and CO in the two peaks using accurate H<sub>2</sub> column densities.

In this study, we present new data from the Five College Radio Astronomy Observatory (FCRAO) 14 m Telescope and the Seoul Radio Astronomy Observatory (SRAO) 6 m Telescope, along with data from the *Spitzer* Space Telescope and Max-Planck Millimetre Bolometer (MAMBO) at the IRAM 30 m Telescope. We combine these data sets to study the reason for the chemical differentiation along the filament. We determine the degree of molecular depletion by comparing column densities calculated from dust continuum emission and molecular line emission, since the dust continuum at (sub) millimeter traces the total material along the line of sight very well. We also carry out chemical calculations to study the abundance evolution of CO and N<sub>2</sub>H<sup>+</sup> with different assumptions about the evolution of the core densities.

Observations of TMC-1 are summarized in Section 2. In Section 3, we show the observational results and simple analysis. Section 4 presents the column densities using several molecules and dust continuum emission. Section 5 describes the chemical modeling sequence and results. Finally, we discuss our results in Section 6.

## 2. Observations

### 2.1. FCRAO Observations

Maps of CS ( $J = 2-1$ ) and N<sub>2</sub>H<sup>+</sup> ( $J = 1-0$ ) toward TMC-1 were made in 2003 January with the 14 m telescope of the FCRAO. The 32-element heterodyne receiver SEQUOIA was used with the autocorrelator configured with a bandwidth of 25 MHz over 1024 channels. The total mapping times were 237 hr and 130 hr for CS ( $J = 2-1$ ) and N<sub>2</sub>H<sup>+</sup> ( $J = 1-0$ ), respectively. The reference center of maps was ( $\alpha$ ,  $\delta$ ) = (4<sup>h</sup>41<sup>m</sup>44<sup>s</sup>.0, +25°42′22″.0).

### 2.2. SRAO Observations

We observed TMC-1 in C<sup>17</sup>O ( $J = 2-1$ ) and C<sup>18</sup>O ( $J = 2-1$ ) with the 6 m telescope of SRAO in 2009 February. The autocorrelator was configured with a bandwidth of 100 MHz over 2048 channels. The pointing uncertainty was less than 10'' in both azimuth and elevation. We observed the positions of the cyanopolyne peak and the ammonia peak in the TMC-1 filament. The total integration time for each position was 30 minutes.

Table 1 summarizes basic observational parameters of the FCRAO and SRAO data, including the frequency ( $\nu$ ), velocity resolution ( $\delta\nu$ ), the FWHM beam size ( $\theta_{\text{mb}}$ ), the main beam efficiency ( $\eta_{\text{mb}}$ ), rms noise level at the given velocity resolution, and the observing date.

**Table 1**  
List of Observed Lines

Molecule	Transition	Telescope	$\nu$ (GHz)	$\delta\nu$ (km s <sup>-1</sup> )	$\theta_{\text{mb}}$ (arcsec)	$\eta_{\text{mb}}$	Rms (K)	Observing Date
N <sub>2</sub> H <sup>+</sup>	$J = 1-0$	FCRAO 14 m	93.174	0.08	58	0.5	0.1	2003 Jan
CS	$J = 2-1$	FCRAO 14 m	97.981	0.07	55	0.5	0.1	2003 Jan
C <sup>18</sup> O	$J = 2-1$	SRAO 6 m	219.560	0.13	48	0.57	0.04	2009 Feb
C <sup>17</sup> O	$J = 2-1$	SRAO 6 m	224.714	0.13	48	0.57	0.04	2009 Feb

### 2.3. The *Spitzer* Space Telescope Observations

The *Spitzer* Legacy Program “From Molecular Cores to Planet Forming Disks” (c2d; Evans et al. 2003) observed TMC-1 at 3.6, 4.5, 5.8, and 8.0  $\mu\text{m}$  with the Infrared Array Camera (IRAC; Fazio et al. 2004) on 2004 October 8 and at 24 and 70  $\mu\text{m}$  with the Multiband Imaging Photometer for *Spitzer* (MIPS; Rieke et al. 2004) on 2004 September 25. The pixel sizes are 1''/2 for all IRAC bands and 2''/45 at MIPS 24  $\mu\text{m}$ . These images were centered at  $(\alpha, \delta) = (4^{\text{h}}41^{\text{m}}14^{\text{s}}, +25^{\circ}59'31'')$  for the 3.6 and 5.8  $\mu\text{m}$  bands and  $(4^{\text{h}}41^{\text{m}}19^{\text{s}}, +25^{\circ}52'49'')$  for the 4.5 and 8.0  $\mu\text{m}$  bands. The MIPS image was centered at  $(\alpha, \delta) = (4^{\text{h}}41^{\text{m}}21^{\text{s}}, +25^{\circ}55'46'')$ . For detailed descriptions of data reduction refer to Evans et al. (2007).

### 2.4. MAMBO Observations

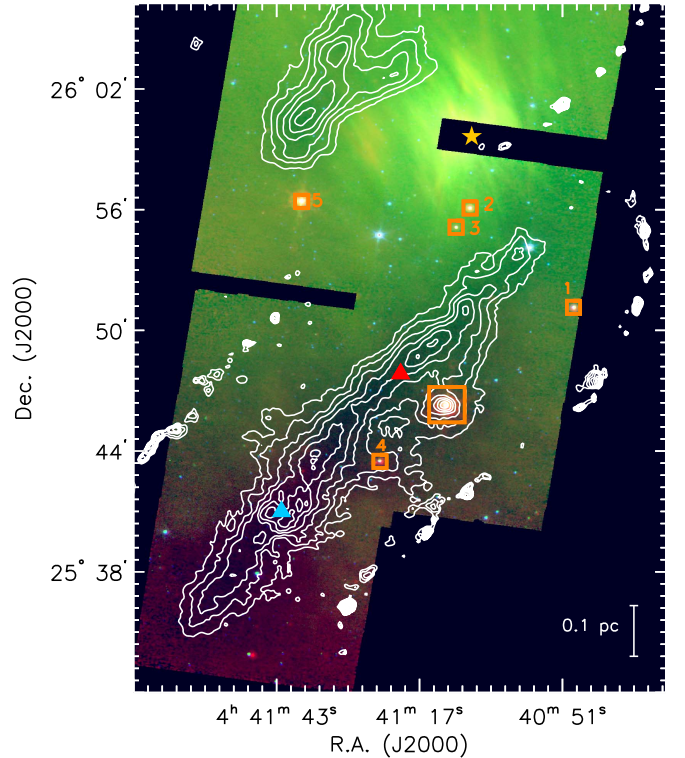
The observations of the 1.2 mm thermal dust continuum emission toward TMC-1 were performed with the IRAM 30 m telescope on Pico Veleta (Spain) using the 117-receiver MAMBO-2 (240'' diameter) camera (Kreysa et al. 1999). The beam size on the sky was 11'', and the effective frequency is 250 GHz with half sensitivity limits at 210 and 290 GHz. TMC-1 was observed between the summer of 2002 and the winter of 2003/2004. The detailed process of the data reduction of MAMBO observations was explained in Kauffmann et al. (2008), where these data were first presented.

## 3. Observational Results and Basic Analysis

We present the relevant data sets in this section, starting with the infrared and submillimeter continuum data, which establish the environmental context for the molecular line observations.

### 3.1. Infrared Observations

Figure 1 shows 1.2 mm dust continuum emission (contours) on top of the *Spitzer* three-color image (4.5, 8.0, and 24  $\mu\text{m}$  for blue, green, and red, respectively). The 1.2 mm dust continuum emission shows that the contours around the cyanopolyyne peak are complex compared to those around the ammonia peak. The IRAS source (IRAS 04381+2540) marked with a large box, is located in the northwest part of TMC-1. The small boxes represent YSO candidates that are classified in the c2d source catalogs (Evans et al. 2007). The source identification numbers are assigned with increasing right ascension. With the *Spitzer* observations, we confirm that there are no YSO candidates along the densest part of the TMC-1 filament, though six YSO candidates (including the IRAS source) are identified outside or at the boundary of the TMC-1 filament. There is a gradient in the 4.5 and 8  $\mu\text{m}$  emission from the northwest to the southeast direction probably due to IRAS 04380+2553, which is a late-B-type star, HD 29647, behind TMC-1 (Whittet et al. 2004; Mooley et al. 2013). The 4.5 and 8  $\mu\text{m}$  radiation from HD 29747 is scattered by dust grains in

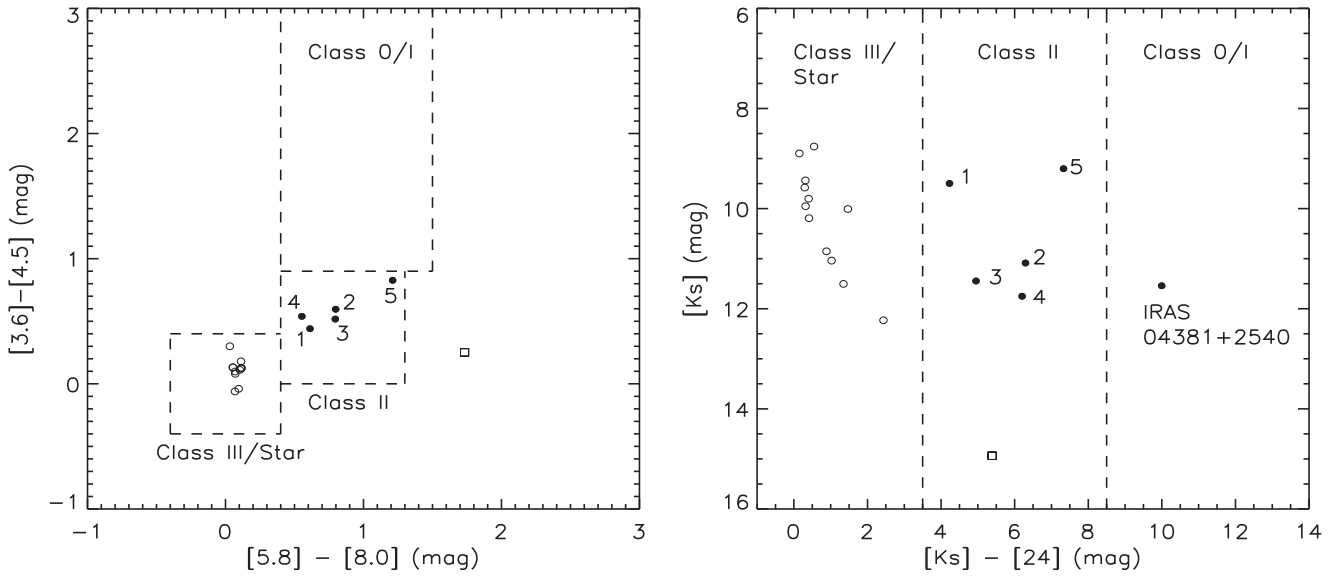


**Figure 1.** MAMBO 1.2 mm dust continuum emission (contours) on top of three-color composite *Spitzer* image of TMC-1. The IRAC 4.5  $\mu\text{m}$ , 8.0  $\mu\text{m}$ , and MIPS 24  $\mu\text{m}$  are presented, respectively, as blue, green, and red. The contours begin at 10 mJy beam<sup>-1</sup> and increase in steps of 5 mJy beam<sup>-1</sup>. Cyanopolyyne peak (blue triangle) and ammonia peak (red triangle) are marked. The large orange box represents IRAS 04381+2540. The small orange boxes indicate classified YSO candidates and they are numbered in order of increasing right ascension. The contours at the top of the image are TMC-1C (Schnee et al. 2007). The yellow star represents the position of HD 29647, but HD 29647 is not shown in this image because the MIPS 24  $\mu\text{m}$  band does not cover the source.

low densities. Therefore, the gradient of the scattered light emission might imply that the TMC-1 filament is inclined toward HD 29647. As a result, the surface of its northwest part is more illuminated by the B star.

To check the evolutionary stages of the six YSO candidates, we use the same color-color and color-magnitude diagrams (Figure 2) as used in Lee et al. (2006). The diagrams use the photometric data from the 2MASS (Two Micron All Sky Survey; Skrutskie et al. 2006), IRAC, and MIPS bands. IRAS 04381+2540 is not included in the color-color diagram since it is not covered by the observations at 3.6 and 5.8  $\mu\text{m}$ . The coordinates and fluxes from the 2MASS and *Spitzer* observations of these YSO candidates are listed in Table 2. As seen in Figure 2, the IRAS source is a Class 0/I candidate, and five YSO candidates are classified as Class II. We compare our





**Figure 2.** Left: color-color diagram for sources in the IRAC bands ([3.6]–[4.5] vs. [5.8]–[8.0]). The filled circles with numbers denote five YSO candidates in TMC-1 (see Figure 1). The open circles and open square indicate stars and a galaxy, respectively, based on the c2d classification (Evans et al. 2007). The dashed boxes indicate the approximate domains of Class 0/I, Class II, and Class III sources as used in Lee et al. (2006). IRAS 04381+2540 is not included because two IRAC bands (3.6 and 5.8  $\mu\text{m}$ ) do not cover the source. Right: color-magnitude diagram for the 2MASS Ks-band and the MIPS 24  $\mu\text{m}$  band sources. This diagram includes IRAS source as well as five YSO candidates. The symbols are the same as in the left diagram. The three domains divided by dashed line show different evolutionary stages.

classification to previous results in the literature. Six YSOs including the IRAS source are in good agreement with previously identified objects in the catalogs of Gutermuth et al. (2008, 2009) and Rebull et al. (2010). Table 3 lists previous studies, our classification, and identifications for six YSOs.

We also calculate the bolometric luminosity ( $L_{\text{bol}}$ ) and bolometric temperature ( $T_{\text{bol}}$ ) for these six YSO candidates using the observed fluxes (see Table 2). The bolometric luminosity ( $L_{\text{bol}}$ ) is calculated by integrating the spectral energy distributions (SEDs), including far-infrared and submillimeter wavelengths by extrapolation

$$L_{\text{bol}} = 4\pi d^2 \int_0^\infty S_\nu d\nu, \quad (1)$$

where  $d$  is the distance to the source and  $S_\nu$  is the flux as a function of frequency of the source. The bolometric temperature ( $T_{\text{bol}}$ ) is described as the temperature of a blackbody with the same flux-weighted mean frequency as the observed SED and it is calculated using the following equation (Myers & Ladd 1993):

$$T_{\text{bol}} = 1.25 \times 10^{-11} \frac{\int_0^\infty \nu S_\nu d\nu}{\int_0^\infty S_\nu d\nu} \text{ K}. \quad (2)$$

Table 4 shows that the evolutionary stage of each source based on the bolometric temperature (Class I :  $70 \text{ K} \leq T_{\text{bol}} \leq 650 \text{ K}$ , Class II :  $650 \text{ K} < T_{\text{bol}} \leq 2880 \text{ K}$ ; Chen et al. 1995) is consistent with the result found using the color-color and color-magnitude diagrams. Furthermore, Dunham et al. (2008) calculated the bolometric luminosity and temperature of  $\sim 0.73 L_\odot$  and  $\sim 165 \text{ K}$  for the IRAS source, which is also classified as a Class I object.

### 3.2. Molecular Line Observations

Figures 3 and 4 present four spectra observed at the cyanopolyne and ammonia peaks. The Gaussian-fitting results of CS 2–1 and  $\text{C}^{18}\text{O}$  2–1 are obtained using the GILDAS/CLASS package<sup>5</sup> and they are summarized in Table 5.  $\text{N}_2\text{H}^+$  1–0 and  $\text{C}^{17}\text{O}$  2–1 have hyperfine structures and they are fitted using the method in the program CLASS (METHOD HFS), which fits the relative intensities of the hyperfine structure lines of a molecular rotational transition to derive the optical depth of the line. The hyperfine-fitting results are listed in Table 6. CS 2–1,  $\text{C}^{18}\text{O}$  2–1, and the isolated component of  $\text{N}_2\text{H}^+$  1–0, especially at the cyanopolyne peak, show double peaks. This double peak feature could be caused either by optical depth or by two velocity components. Here, we assume that the line shape is caused by optical depth, rather than two velocity components, and fit the data with a single Gaussian.

Figures 5 and 6 present the integrated intensity maps of CS 2–1 (green line) and  $\text{N}_2\text{H}^+$  1–0 (yellow line) obtained at the FCRAO. In the figure, line emissions (contours) are overlaid on the 1.2 mm dust continuum emission map (gray image), and the locations of the cyanopolyne peak and the ammonia peak are denoted by red and blue triangles, respectively. This 1.2 mm dust continuum emission map, the best tracer of column density, is the first 1.2 mm dust continuum emission map of its type for the whole TMC-1 region (Kauffmann et al. 2008). The CS 2–1 intensity does not peak at the cyanopolyne peak, while it peaks close to the ammonia peak. However, the  $\text{N}_2\text{H}^+$  1–0 intensity peaks around both cyanopolyne and ammonia peaks, though the intensity close to the cyanopolyne peak is weaker than that around the ammonia peak. The differences of the molecular line intensities at the two peaks can be caused by differences in (1) density, (2) temperature, (3) abundance, or any combination of these properties.

<sup>5</sup> <http://www.iram.fr/IRAMFR/GILDAS/>

**Table 2**  
Fluxes of Young Stellar Object Candidates in TMC-1

Source <sup>a</sup>	c2d ID	R.A. (J2000) ( <sup>h</sup> <sup>m</sup> <sup>s</sup> )	Decl. (J2000) ( <sup>°</sup> <sup>'</sup> <sup>''</sup> )	Fluxes from 2MASS (mJy)			Fluxes from <i>Spitzer</i> (mJy)					
				J (1.235 $\mu$ m)	H (1.662 $\mu$ m)	Ks (2.159 $\mu$ m)	IRAC1 (3.6 $\mu$ m)	IRAC2 (4.5 $\mu$ m)	IRAC3 (5.8 $\mu$ m)	IRAC4 (8.0 $\mu$ m)	MIPS1 (24 $\mu$ m)	MIPS2 (70 $\mu$ m)
YSO 1	SSTc2d J044049.5+255119	04 40 49.5	25 51 19.1	7.99E+01	1.10E+02	1.06E+02	6.69E+01	6.42E+01	5.54E+01	5.43E+01	5.58E+01	...
YSO 2	SSTc2d J044108.2+255607	04 41 08.2	25 56 07.4	5.11E+00	1.42E+01	2.45E+01	3.71E+01	4.11E+01	4.46E+01	5.19E+01	8.64E+01	...
YSO 3	SSTc2d J044110.8+255512	04 41 10.8	25 55 11.6	8.42E+00	1.45E+01	1.76E+01	1.32E+01	1.36E+01	1.37E+01	1.59E+01	1.80E+01	...
IRAS source <sup>b</sup>	SSTc2d J044112.6+254635	04 41 12.6	25 46 35.4	2.20E−01	2.27E+00	1.61E+01	...	1.39E+02	...	3.46E+02	1.72E+03	6.74E+03
YSO 4	SSTc2d J044124.6+254353	04 41 24.6	25 43 53.0	3.30E−01	3.95E+00	1.33E+01	1.95E+01	2.05E+01	2.09E+01	1.94E+01	4.31E+01	...
YSO 5	SSTc2d J044138.8+255627	04 41 38.8	25 56 26.8	2.90E+01	9.14E+01	1.39E+02	1.16E+02	1.59E+02	2.48E+02	4.22E+02	1.27E+03	2.11E+03

**Notes.**

<sup>a</sup> The numbers refer to the marked YSOs in Figure 1.

<sup>b</sup> IRAS 04381+2540.

**Table 3**  
Comparison with Previous Studies

Source <sup>a</sup>	c2d ID	Gutermuth Category <sup>b</sup>	Rebull Category <sup>c</sup>	This Study	Identification <sup>d</sup>
YSO 1	SSTc2d J044049.5+255119	Class II	Class II	II	JH 223
YSO 2	SSTc2d J044108.2+255607	Class II	Flat	II	ITG 33A
YSO 3	SSTc2d J044110.8+255512	Class II	Class II	II	ITG 34
IRAS source	SSTc2d J044112.6+254635	Embedded	Class I	0/I	IRAS 04381+2540
YSO 4	SSTc2d J044124.6+254353	Class II	Flat	II	GKH 32
YSO 5	SSTc2d J044138.8+255627	Class II	Flat	II	Haro 6-33

**Notes.**

<sup>a</sup> The numbers refer to the marked YSOs in Figure 1.

<sup>b</sup> Gutermuth et al. (2008, 2009).

<sup>c</sup> Rebull et al. (2010).

<sup>d</sup> Reference for the first identification of each source: JH 223—Jones & Herbig (1979), ITG 33A, ITG 34—Itoh et al. (1999), IRAS 04381+2540—Beichman et al. (1986), GKH 32—Gomez et al. (1994), Haro 6-33—Haro et al. (1953).

**Table 4**  
Results of Bolometric Luminosity and Temperature of Six YSO Candidates in TMC-1

Source <sup>a</sup>	c2d ID	$L_{\text{bol}} (L_{\odot})$	$T_{\text{bol}} (\text{K})$	Class
YSO 1	SSTc2d J044049.5+255119	0.12	1722	II
YSO 2	SSTc2d J044108.2+255607	0.04	991	II
YSO 3	SSTc2d J044110.8+255512	0.02	1460	II
IRAS source <sup>b</sup>	SSTc2d J044112.6+254635	0.54	214	I
YSO 4	SSTc2d J044124.6+254353	0.02	922	II
YSO 5	SSTc2d J044138.8+255627	0.33	758	II

**Notes.**

<sup>a</sup> The numbers refer to the marked YSOs in Figure 1.

<sup>b</sup> IRAS 04381+2540.

We use the  $\text{N}_2\text{H}^+$  1–0 line to study the velocity field because that line has the lowest optical depth. Figure 7 presents the first moment map. We found three velocity components; the northern part (red to yellow), the southern part (green), and the northwestern part (sky blue). In the northern part (P1) of the TMC-1 filament, the velocity is  $\sim 5.8 \text{ km s}^{-1}$  and it drops to about  $5.4 \text{ km s}^{-1}$  in the southern part (P5). On the other hand, the source velocity in the northwestern part (IRAS 04381+2540,  $V_{\text{LSR}} \sim 5.2 \text{ km s}^{-1}$ ) is lower than the typical velocity along the TMC-1 filament ( $V_{\text{LSR}} \sim 5.6 \text{ km s}^{-1}$ ), suggesting that this region may be separated from the TMC-1 filament along the line of sight (Toelle et al. 1981; Snell et al. 1982; Kolotovkina et al. 1986; Olano et al. 1988; Fehér et al. 2016).

Table 7 summarizes the source velocities of five positions (P1, P2, P3, P4, and P5) along the TMC-1 filament, consistent with the MAMBO dust continuum peaks (Kauffmann et al. 2008) and the IRAS source (IRAS 04381+2540).

#### 4. Column Density Calculations

##### 4.1. $\text{H}_2$ Column Density from Dust Continuum Emission

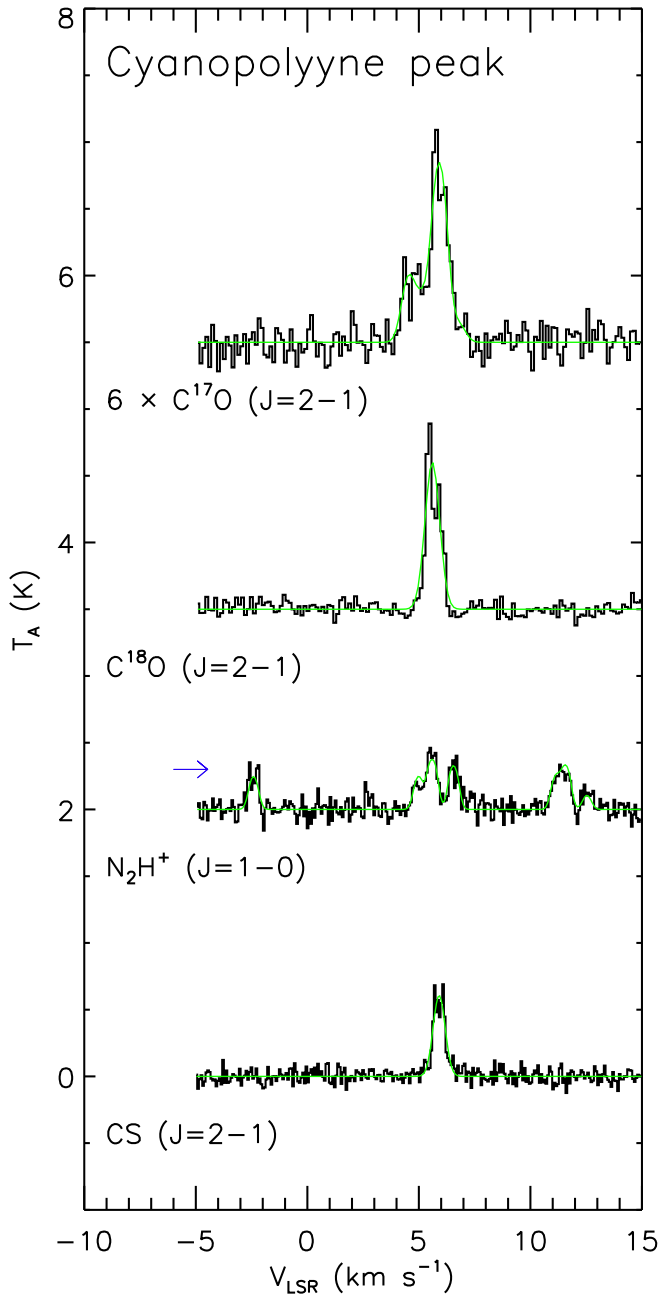
We compare  $\text{H}_2$  column densities calculated from molecular line emission and dust continuum emission to calculate the degree of depletion of molecules.

If the dust continuum emission is optically thin, the observed flux  $F_{\nu}^{\text{beam}}$  can be related to the column density of gas. Therefore, we calculate the  $\text{H}_2$  column density with the MAMBO flux as given by this equation (Hildebrand 1983):

$$N(\text{H}_2) = \frac{F_{\nu}^{\text{beam}}}{\Omega_{\text{A}} \mu_{\text{H}_2} m_{\text{H}} \kappa_{\nu} B_{\nu}(T)}, \quad (3)$$

where  $F_{\nu}^{\text{beam}}$  is the observed flux per beam,  $\mu_{\text{H}_2}$  ( $=2.3$ ) is the molecular weight per hydrogen molecule,  $m_{\text{H}}$  is the atomic mass unit,  $\kappa_{\nu}$  ( $=0.0102 \text{ cm}^2 \text{ g}^{-1}$ ) is the mass opacity of dust per gram of gas at 1.2 mm (Kauffmann et al. 2008),  $B_{\nu}(T)$  is the Planck function, and  $\Omega_{\text{A}}$  is the beam solid angle;  $\Omega_{\text{A}} = (\pi \theta_{\text{HPBW}}^2) / (4 \ln 2)$  for a Gaussian function (Lee et al. 2003; Kauffmann et al. 2008). The MAMBO map is convolved with a Gaussian profile with the FWHM of the beam of molecular observations. As a result, the  $\text{H}_2$  column density calculated from dust continuum emission at the cyanopolyne and the ammonia peak is  $(1.99 \pm 0.20) \times 10^{22} \text{ cm}^{-2}$  and  $(1.87 \pm 0.19) \times 10^{22} \text{ cm}^{-2}$ , respectively. The  $\text{H}_2$  column densities calculated from the dust continuum emission, the best tracer of the column density along the line of sight, is essentially identical at the two peaks.

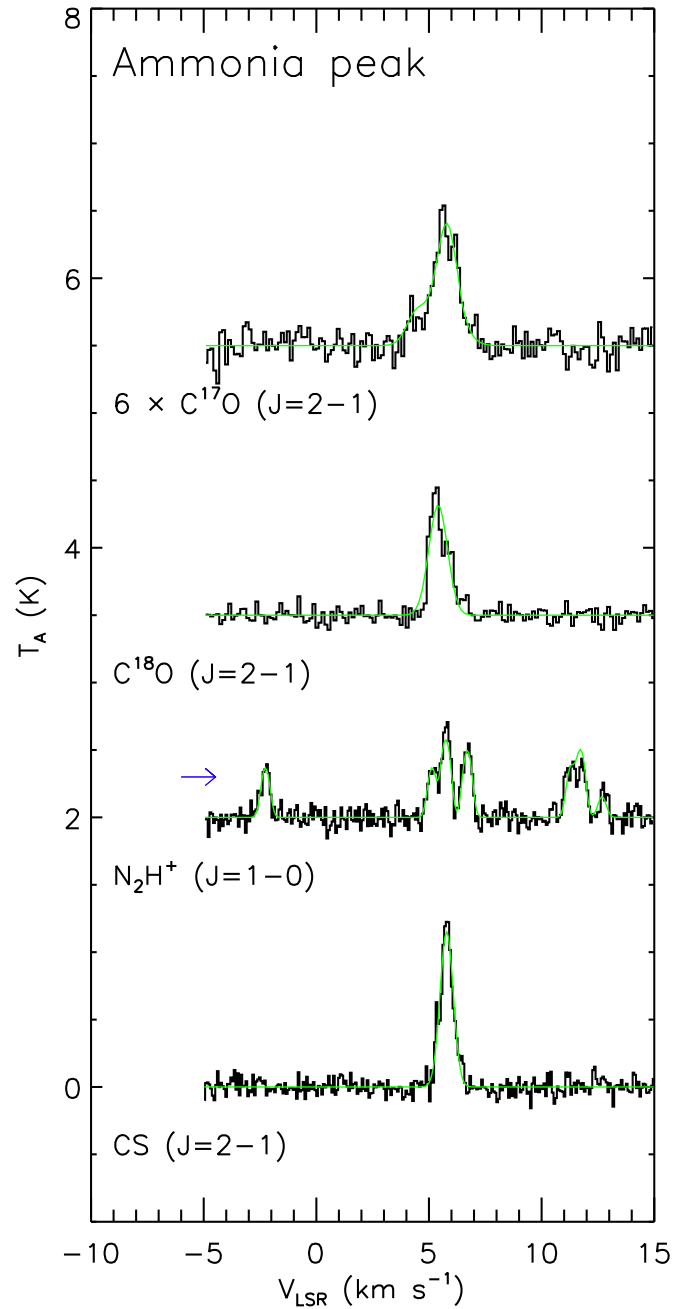
Two other recent studies also concluded that the column densities in the ammonia and cyanopolyne peaks are comparable. Suutarinen et al. (2011) estimated the total  $\text{H}_2$  column density in the TMC-1 region using the SCUBA 850 and 450  $\mu\text{m}$  data from Nutter et al. (2008). The derived  $N(\text{H}_2)$  are  $\sim 1.3 \times 10^{22} \text{ cm}^{-2}$  and  $\sim 1.6 \times 10^{22} \text{ cm}^{-2}$  at the ammonia peak and the cyanopolyne peak, respectively. In addition, Fehér et al. (2016) used the *Herschel*/SPIRE images to derive the column densities of  $3.3 \times 10^{22} \text{ cm}^{-2}$  at the ammonia peak and  $2.7 \times 10^{22} \text{ cm}^{-2}$  at the cyanopolyne peak. These absolute values are different from ours probably because of different beam sizes. However, the very important point of these calculations including ours is that the column densities at the ammonia and cyanopolyne peaks are similar, contrary to suggestions by some previous studies (see Section 1).



**Figure 3.** Spectra of molecular transitions observed with the FCRAO 14 m and the SRAO 6 m telescopes at the cyanopolyne peak (black). The Gaussian-fitting results and hyperfine-fitting results for CS 2-1/ $C^{18}O$  2-1 and  $N_2H^+$  1-0/ $C^{17}O$  2-1 are shown as green lines, respectively. The isolated component of  $N_2H^+$  1-0 is pointed out by the blue arrow.

#### 4.2. Column Density and Abundance for Molecular Line Emission

Chemical models predict that CO and  $N_2H^+$  will be anti-correlated because CO destroys  $N_2H^+$  in the gas phase. Previous observations of pre- and protostellar regions (Bergin et al. 2001; Tafalla et al. 2002; Lee et al. 2003; Di Francesco et al. 2004; Jørgensen et al. 2004; Pagani et al. 2005) confirmed that  $N_2H^+$  becomes abundant as CO is frozen onto grain surfaces. Therefore, these two molecules together can trace different chemical evolutionary stages of molecular cores. We



**Figure 4.** Spectra of molecular transitions observed with the FCRAO 14 m and the SRAO 6 m telescopes at the ammonia peak. The Gaussian-fitting results and hyperfine-fitting results for CS 2-1/ $C^{18}O$  2-1 and  $N_2H^+$  1-0/ $C^{17}O$  2-1 are shown as green lines, respectively. The isolated component of  $N_2H^+$  1-0 is pointed out by the blue arrow.

use the  $N_2H^+$  and  $C^{17}O$  lines to calculate column densities more accurately because these lines have hyperfine structures. In this calculation, we assume that (1) the dust continuum emission traces all material along a line of sight; (2) dust and gas are well coupled thermally ( $T_d = T_k = 10$  K, Pratap et al. 1997); (3) molecular abundances are constant along the line of sight; and (4) all levels are in LTE, following Lee et al. (2003).

For line emission, if a line is optically thin ( $\tau < 0$ ) for the transition  $J = J \rightarrow J - 1$  in a linear molecule, the relation

**Table 5**  
Results of Gaussian Fits to CS ( $J = 2-1$ ) and C<sup>18</sup>O ( $J = 2-1$ ) Lines

Line	Cyanopolyne Peak				Ammonia Peak			
	$V_{\text{LSR}}$ (km s <sup>-1</sup> )	$\Delta V$ (km s <sup>-1</sup> )	$T_A$ (K)	$\int T_A dV$ (K km s <sup>-1</sup> )	$V_{\text{LSR}}$ (km s <sup>-1</sup> )	$\Delta V$ (km s <sup>-1</sup> )	$T_A$ (K)	$\int T_A dV$ (K km s <sup>-1</sup> )
CS ( $J = 2-1$ )	5.91 (0.01)	0.63 (0.03)	0.60 ...	0.40 (0.01)	5.78 (0.01)	0.70 (0.02)	1.16 ...	0.86 (0.02)
C <sup>18</sup> O ( $J = 2-1$ )	5.61 (0.01)	0.76 (0.02)	1.08 ...	0.87 (0.01)	5.42 (0.02)	0.94 (0.05)	0.80 ...	0.80 (0.03)

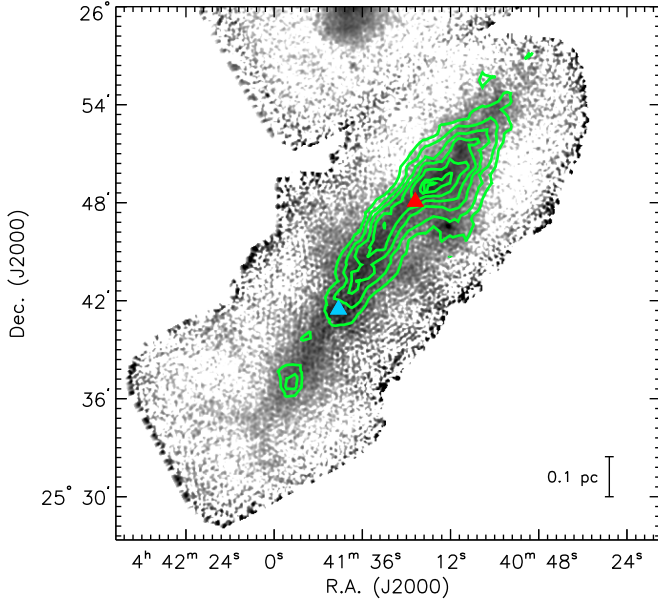
**Note.** The values in parentheses refer to errors of each parameter.

**Table 6**  
Results of Hyperfine Structure to N<sub>2</sub>H<sup>+</sup> ( $J = 1-0$ ) and C<sup>17</sup>O ( $J = 2-1$ ) Lines

Line	Cyanopolyne Peak			Ammonia Peak		
	$V_{\text{LSR}}$ (km s <sup>-1</sup> )	$\Delta V$ (km s <sup>-1</sup> )	$\tau^a$	$V_{\text{LSR}}$ (km s <sup>-1</sup> )	$\Delta V$ (km s <sup>-1</sup> )	$\tau^a$
N <sub>2</sub> H <sup>+</sup> ( $J = 1-0$ )	5.60 (0.01)	0.44 (0.03)	7.64 (2.11)	5.75 (0.01)	0.41 (0.02)	6.77 (1.63)
C <sup>17</sup> O ( $J = 2-1$ )	5.65 (0.02)	0.68 (0.05)	1.71 (0.66)	5.53 (0.02)	1.01 (0.04)	0.10 (0.16)

**Note.** The values in parentheses refer to errors of each parameter.

<sup>a</sup>  $\tau$  is the sum of the optical depths of all hyperfine components.



**Figure 5.** Integrated intensity maps of CS 2–1 (green contours) on top of the 1200  $\mu\text{m}$  dust continuum emission map (gray scale). Blue and red triangles denote the cyanopolyne peak and the ammonia peak, respectively. The contour values are 0.38, 0.53, 0.68, 0.83, 0.98, 1.13, and 1.17 K km s<sup>-1</sup>.

between the column density  $N(x)$  of molecule  $x$  and the integrated intensity of the line is

$$N(x) = \frac{3kQe^{E_J/kT_{\text{ex}}}}{8\pi^3\nu\mu^2J} \int T_R dv \quad (4)$$

where  $T_{\text{ex}}$  is the excitation temperature above the ground state, and  $\mu$  is the dipole moment. The frequency ( $\nu$ ), energy above ground ( $E_J$ ), and the partition function ( $Q$ ) can be computed from  $B$ , the rotational constant.

However, if the optical depth is not negligible, the optical depth effect is corrected for with the equation

$$N_{\text{thick}} = N_{\text{thin}} \frac{\tau_\nu}{1 - e^{-\tau_\nu}}. \quad (5)$$

Additionally, we calculate the abundance of molecule,  $X(x)$ , using the H<sub>2</sub> column density calculated from the dust continuum emission

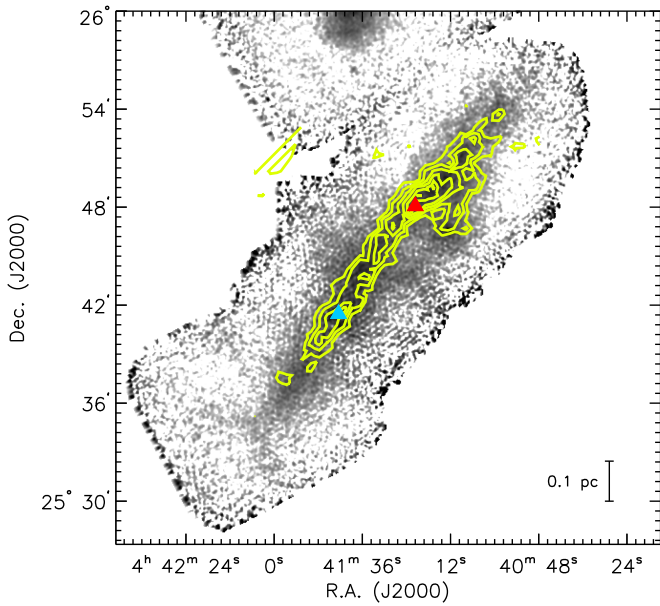
$$X(x) = \frac{N(x)}{N(\text{H}_2)_{\text{Dust}}}. \quad (6)$$

Table 8 summarizes the parameters (the rotational constant  $B$  and the dipole moment  $\mu$ ) used for the calculation of the column densities. The optical depths of the N<sub>2</sub>H<sup>+</sup> 1–0 and C<sup>17</sup>O 2–1 lines are obtained by the hyperfine-fitting method in the CLASS package (see Table 6). We find that C<sup>17</sup>O 2–1 is optically thick at the cyanopolyne peak, but optically thin at the ammonia peak. On the other hand, the total optical depth of N<sub>2</sub>H<sup>+</sup> is too large to adopt Equations (4) and (5). Therefore, we use the isolated component (see Figures 3 and 4) to derived the column density of N<sub>2</sub>H<sup>+</sup> by assuming that the optical depth of each component is proportional to its LTE intensity ratio. The optical depth of the isolated component is 10% of the total optical depth.

The column densities and abundances for molecular lines are listed in Table 9. The column density from C<sup>17</sup>O 2–1 is a factor of two larger at the cyanopolyne peak than at the ammonia peak. However, there is no significant difference in the column densities of N<sub>2</sub>H<sup>+</sup> 1–0 between the cyanopolyne and the ammonia peaks when the error is considered.

Table 10 represents the depletion factors of C<sup>17</sup>O and N<sub>2</sub>H<sup>+</sup> when we compare the calculated abundances with the canonical abundances,  $X(\text{C}^{17}\text{O}) = 1.5 \times 10^{-7}$  and  $X(\text{N}_2\text{H}^+) = 5.0 \times 10^{-10}$ . This canonical abundance of C<sup>17</sup>O is obtained by assuming <sup>16</sup>O/<sup>18</sup>O of 540, <sup>18</sup>O/<sup>17</sup>O of 3.65 (Penzias 1981; Wilson & Rood 1994), and CO/H<sub>2</sub> of  $2.7 \times 10^{-4}$  (Lacy et al. 1994). The





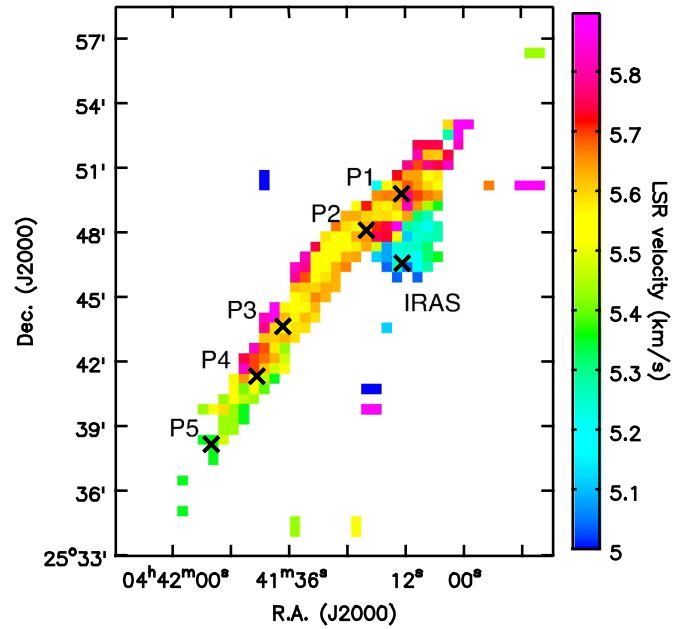
**Figure 6.** Integrated intensity maps of  $\text{N}_2\text{H}^+$  1–0 (yellow contours) on top of the  $1200\ \mu\text{m}$  dust continuum emission map (gray scale). Blue and red triangles denote the cyanopolyne peak and the ammonia peak, respectively. The contour values are 0.08, 0.12, 0.16, 0.20, 0.24, and  $0.27\ \text{K km s}^{-1}$ .

canonical abundance of  $\text{N}_2\text{H}^+$  is adopted from Johnstone et al. (2010). We find that the depletion of CO is more significant in the ammonia peak (CO depletion factor = 11.9) compared to the cyanopolyne peak (CO depletion factor = 4.5). This suggests that the ammonia peak is more chemically evolved than the cyanopolyne peak. On the other hand, the depletion of  $\text{N}_2\text{H}^+$  is almost identical in the two peaks if errors are considered.

### 5. Chemical Modeling

We calculate the evolution of chemical abundances of CO and  $\text{N}_2\text{H}^+$  molecules to explain the differential chemical distribution along the TMC-1 filament. According to our analysis in the previous section, the cyanopolyne and ammonia peaks have similar column densities calculated from the dust continuum emission, while CO is depleted more significantly at the ammonia peak than at the cyanopolyne peak. This trend implies that longer timescales at lower densities are necessary at the ammonia peak during the density evolution since in that condition,  $\text{N}_2$  forms more in the gas and CO is frozen-out more onto grain surfaces, resulting in a high abundance of  $\text{N}_2\text{H}^+$  with a significant depletion of CO. Therefore, we calculate the chemical evolution by varying the timescales over which density evolves to account for the observational results.

We adopt the chemical network used in Lee et al. (2004) and the updated  $\text{N}_2$  binding energy used in Chen et al. (2009). In the chemical network, the interactions between gas and dust grains (i.e., freeze-out and evaporation of molecules on and off grain surfaces) are included as well as the pure gas chemistry. We assume that the grain surface is coated by water ice and use the same initial elemental abundances as presented in Table 3 of Lee et al. (2004). According to previous studies (Hirota et al. 1998, 2003), the two peaks have a current density of  $\sim 10^5\ \text{cm}^{-3}$ . We therefore consider that the densities of the two peaks evolve from  $10^3$  to  $10^5\ \text{cm}^{-3}$  and assume that a core is formed when a density condensation is larger than  $10^3\ \text{cm}^{-3}$ ,



**Figure 7.** First moment (i.e., intensity weighted central velocity) map of the  $\text{N}_2\text{H}^+$  1–0 line. Five positions (P1, P2, P3, P4, and P5) are placed along the TMC-1 filament in order of increasing right ascension. The IRAS source (IRAS 04381+2540) is shown in the northwest part of the TMC-1 filament. The coordinate and  $V_{\text{LSR}}$  of each position are listed in Table 7.

**Table 7**  
LSR Velocity in the TMC-1 Filament

Position	R.A. (J2000) ( $^{\text{h}}\ ^{\text{m}}\ ^{\text{s}}$ )	Decl. (J2000) ( $^{\circ}\ ' \ ''$ )	LSR Velocity ( $\text{km s}^{-1}$ )
The Direction Along the TMC-1 Filament			
P1	04 41 13.8	25 49 49.5	5.82 (0.01)
P2 <sup>a</sup>	04 41 20.9	25 48 07.2	5.75 (0.01)
P3	04 41 37.2	25 43 38.5	5.66 (0.01)
P4 <sup>b</sup>	04 41 42.5	25 41 27.0	5.60 (0.01)
P5	04 41 51.9	25 38 09.5	5.42 (0.01)
IRAS <sup>c</sup>	04 41 12.8	25 46 33.5	5.20 (0.01)

**Notes.** The values in parentheses refer to errors of LSR velocities.

<sup>a</sup> Ammonia peak.

<sup>b</sup> Cyanopolyne peak.

<sup>c</sup> IRAS 04381+2540.

**Table 8**  
Parameters for Molecules

Line	Rotational Constant B (MHz)	Dipole Moment $\mu$ (D <sup>a</sup> )
$\text{N}_2\text{H}^+$ 1–0	46586.87	3.40
$\text{C}^{17}\text{O}$ 2–1	56179.99	0.11

**Notes.** Molecular spectroscopy data are from the CDMS catalog (Müller et al. 2005, <http://www.astro.uni-koeln.de/cdms>).

<sup>a</sup> The Debye (D) is a unit of the electric dipole moment,  $1\ \text{D} = 10^{-18}\ \text{g}^{1/2}\ \text{cm}^{5/2}\ \text{s}$ .

which is the lowest density at which prestellar cores can be identified (Ward-Thompson et al. 2007). Table 11 lists the timescale of density evolution, i.e., the amount of time spent from

**Table 9**  
Column Densities and Abundances for Molecular Lines

Line	Cyanopolyne Peak		Ammonia Peak	
	Column Density ( $\times 10^{13} \text{ cm}^{-2}$ )	Abundance ( $\times 10^{-9}$ )	Column Density ( $\times 10^{13} \text{ cm}^{-2}$ )	Abundance ( $\times 10^{-9}$ )
$\text{N}_2\text{H}^+$ 1–0	0.5 (0.2)	0.3 (0.1)	0.6 (0.2)	0.3 (0.1)
$\text{C}^{17}\text{O}$ 2–1	66.5 (26.7)	33.4 (13.9)	25.1 (3.1)	13.4 (2.0)

**Note.** We correct for the optical depth effect to calculate the column density from  $\text{N}_2\text{H}^+$  1–0 and  $\text{C}^{17}\text{O}$  2–1. The values in parentheses refer to errors of each parameter.

**Table 10**  
Depletion Factors of CO and  $\text{N}_2\text{H}^+$

	Depletion Factors ( $X_{\text{canonical}}/X_{\text{measured}}$ )	
	Cyanopolyne Peak	Ammonia Peak
$\text{C}^{17}\text{O}$	4.5 (1.9)	11.9 (1.9)
$\text{N}_2\text{H}^+$	1.8 (0.7)	1.6 (0.6)

**Note.** The values in parentheses refer to errors of depletion factors.

**Table 11**  
Timescale of Density Evolution for Chemical Modeling

Gas Density <sup>b</sup> ( $\text{cm}^{-3}$ )	Timescale <sup>a</sup> (years)	
	Cyanopolyne Peak Model	Ammonia Peak Model
$10^3$	$1.0 \times 10^6$	$1.5 \times 10^6$
$10^4$	$1.0 \times 10^5$	$6.0 \times 10^5$
$3 \times 10^4$	$4.0 \times 10^4$	$4.0 \times 10^4$

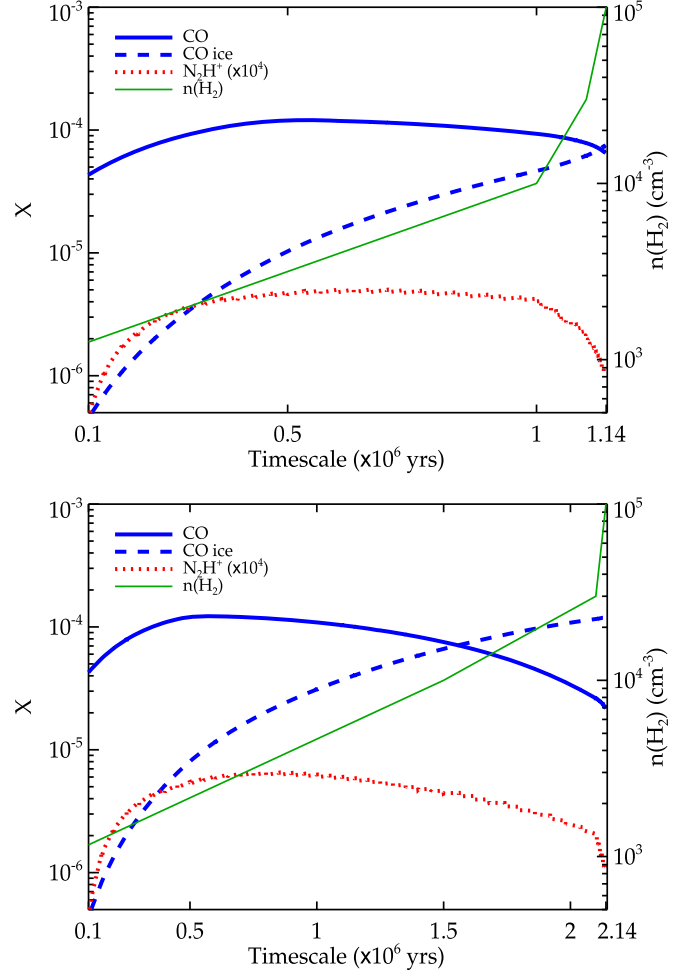
**Notes.**

<sup>a</sup> Timescale spent from a given density to the next density. The final density is  $10^5 \text{ cm}^{-3}$ .

<sup>b</sup> Gas density with the mean molecular weight of  $\mu = 2.3$ .

each density to the next density until the density of the two peaks reaches  $10^5 \text{ cm}^{-3}$  in the models, which fit the observations the best. The densities in the two models gradually increase as time goes by (see Figure 8). The timescale spent at low densities (until the density reaches  $3 \times 10^4 \text{ cm}^{-3}$ ) is shorter in the cyanopolyne peak model ( $1.1 \times 10^6$  years) than in the ammonia peak model ( $2.1 \times 10^6$  years); the density evolves faster in the cyanopolyne peak model than the ammonia peak model until the density reaches  $3 \times 10^4 \text{ cm}^{-3}$ .

Figure 8 presents the abundance evolution of CO and  $\text{N}_2\text{H}^+$  in the best-fit models for the cyanopolyne peak and the ammonia peak, and the abundances of the two molecules at  $n(\text{H}_2) = 10^5 \text{ cm}^{-3}$  in the two models are listed in Table 12. During the density evolution, the CO freeze-out occurs in the two models, but it is more significant in the ammonia peak model than the cyanopolyne peak model, as seen in the observations (see Table 10). The abundance of CO in the cyanopolyne peak model is higher than that of CO in the ammonia peak model by a factor of  $\sim 3$  at  $n(\text{H}_2) = 10^5 \text{ cm}^{-3}$ . On the other hand, the  $\text{N}_2\text{H}^+$  abundances in the two models are similar at  $n(\text{H}_2) = 10^5 \text{ cm}^{-3}$ ; the  $\text{N}_2\text{H}^+$  abundance decreases quickly at  $n(\text{H}_2) = 10^4 \text{ cm}^{-3}$  in the cyanopolyne peak model because of its destruction via abundant CO, while the  $\text{N}_2\text{H}^+$  abundance declines slowly in the ammonia peak model because its destroyer, CO, is already depleted at the low density. In the cyanopolyne peak model, the maximum abundance of  $\text{N}_2\text{H}^+$  is lower than that in the ammonia peak model because its



**Figure 8.** Abundance evolution of CO and  $\text{N}_2\text{H}^+$  in the cyanopolyne peak model (upper) and the ammonia peak model (lower) from the chemical calculations. The blue solid and dotted lines represent CO abundances in gas and ice, respectively. The red dotted line indicates the  $\text{N}_2\text{H}^+$  abundance and the green solid line represents gas density. The timescales until the density reaches  $10^5 \text{ cm}^{-3}$  are  $1.14 \times 10^6$  years in the cyanopolyne peak model and  $2.14 \times 10^6$  years in the ammonia peak model; the maximum timescales in the horizontal axis differ between the two panels.

destruction by CO is more efficient in the cyanopolyne peak model.

According to our modeling, the timescale difference of  $1 \times 10^6$  years in density evolution is needed in order to explain the chemical differentiation between the cyanopolyne and ammonia peaks; the cyanopolyne peak has a shorter timescale than the ammonia peak. This implies that the core at the ammonia peak forms earlier than the core at the cyanopolyne peak, which is consistent with the suggestion by previous

**Table 12**  
Abundances of CO and  $\text{N}_2\text{H}^+$  at  $n(\text{H}_2) = 10^5 \text{ cm}^{-3}$  in the Chemical Calculation

Molecule	Abundance	
	Cyanopolyne Peak Model	Ammonia Peak Model
CO <sup>a</sup>	$6.50 \times 10^{-5}$	$2.17 \times 10^{-5}$
$\text{N}_2\text{H}^+$ <sup>b</sup>	$1.06 \times 10^{-10}$	$1.03 \times 10^{-10}$

**Notes.**

<sup>a</sup> The maximum CO abundance is  $1.2 \times 10^{-4}$  in both models.

<sup>b</sup> The maximum abundance of  $\text{N}_2\text{H}^+$  is different in the two models ( $5 \times 10^{-10}$  in the cyanopolyne peak model and  $6.5 \times 10^{-10}$  in the ammonia peak model).

studies; core formation propagated from the ammonia peak to the cyanopolyne peak (Hanawa et al. 1994; Howe et al. 1996; Markwick et al. 2000, 2001; Saito et al. 2002). However, the current densities of the two peaks are similar. Therefore, once the core at the ammonia peak forms, it might evolve very slowly until the core at the cyanopolyne peak forms and catches up in its density evolution. According to our chemical models combined with the density evolution of cores, the more evolved chemistry at the ammonia peak can be attributed to the longer timescale of density evolution at the peak.

Lee et al. (2003) suggested that the chemical evolution of a core depends on its dynamical timescale as well as its density structure by comparing the chemical distributions in three isolated prestellar cores (L1512, L1544, and L1689B). According to Lee et al. (2003), L1544 and L1689B have the same density distribution, but L1689B is chemically much younger than L1544. In conclusion, L1689B might experience a free-fall-like dynamical process, while L1544 might undergo an ambipolar diffusion process for the dynamical evolution. In the TMC-1 filament, the core associated with the ammonia peak requires a longer evolutionary timescale, so it would have formed via ambipolar diffusion. On the other hand, the core associated with the cyanopolyne peak might have formed by a more violent dynamical process. Some evidence for the dynamical process on a short timescale is present in our observations. First, the 1.2 mm dust continuum emission, which traces the total amount of material the best, shows a complex structure around the cyanopolyne peak (see Figure 1). Second, the central velocity of the  $\text{N}_2\text{H}^+$  line changes sharply around the cyanopolyne peak (see Figure 7). Finally, the molecular line profiles, especially the isolated component of  $\text{N}_2\text{H}^+ J = 1-0$ , show a double peak feature (see Figure 3), indicative of two velocity components. (CS 2–1 and  $\text{C}^{18}\text{O}$  2–1 are also affected by optical depths.) From this evidence, we could consider that two molecular clouds are colliding around the cyanopolyne peak to induce a fast core formation.

## 6. Summary

We study a possible interpretation of the differentiated chemical distribution in the TMC-1 region using various observational data and chemical modeling.

From *Spitzer* infrared observations, we confirm that there are no YSO candidates along the TMC-1 filament, though there are six YSO candidates located outside or at boundary of the TMC-1 filament; five classified YSO candidates objects are classified as Class II sources, and IRAS source (IRAS 04381+2540) is a Class I

object, which agree with previous classifications (Dunham et al. 2008; Gutermuth et al. 2008, 2009; Rebull et al. 2010).

The CS 2–1 and  $\text{N}_2\text{H}^+$  1–0 lines are stronger at the ammonia peak than at the cyanopolyne peak. On the other hand,  $\text{C}^{17}\text{O}$  2–1 and  $\text{C}^{18}\text{O}$  2–1 lines are stronger at the cyanopolyne peak than at the ammonia peak.

Using the  $\text{C}^{17}\text{O}$  and  $\text{N}_2\text{H}^+$  lines, which have hyperfine structures, we calculate the  $\text{H}_2$  column density to compare with the column density calculated from the dust continuum emission, which traces the material the best along the line of sight. The column densities calculated from the dust continuum emission and from the  $\text{N}_2\text{H}^+$  1–0 line are similar at the cyanopolyne and ammonia peaks, while the column density calculated from  $\text{C}^{17}\text{O}$  2–1 is much greater at the cyanopolyne peak than at the ammonia peak.

Comparing the depletion factors of CO and  $\text{N}_2\text{H}^+$  (see Table 10), we find that the CO molecule is much less depleted at the cyanopolyne peak than at the ammonia peak, while the  $\text{N}_2\text{H}^+$  abundance is similar at both peaks, ruling out the possibility of a late-time secondary abundance peak of carbon-chain molecules (Ruffle et al. 1997, 1999; Li et al. 2002; Lee et al. 2003) because both CO and  $\text{N}_2\text{H}^+$  are significantly depleted in the late stage of chemical evolution. Therefore, the relative depletion factors of CO and  $\text{N}_2\text{H}^+$  imply that the cyanopolyne peak is chemically younger (see Section 1).

According to our chemical modeling, the ammonia peak stays for a longer time at low density ( $10^3$ – $10^5 \text{ cm}^{-3}$ ) than the cyanopolyne peak to explain the chemical difference between the two peaks, indicating that the dynamical evolution can affect the chemical evolution, resulting in different chemical states even at the same final density. As a result, we suggest that the chemical differentiation along the TMC-1 filament can be induced by different dynamical processes of core formation; the ambipolar diffusion for the ammonia peak and a free-fall like process for the cyanopolyne peak. Some evidence, such as the complex distribution of the 1.2 mm continuum emission (see Figure 1), the sharp shift of the central velocity of  $\text{N}_2\text{H}^+$  1–0 (see Figure 7), and the double peak feature of molecular line profiles around the cyanopolyne peak (see Figure 3) suggest that two molecular clouds are colliding around the cyanopolyne peak to induce a fast core formation.

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