

Preliminary Data Access Center : User Report

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ABSTRACT

A report on user experience of the Preliminary Data Access Center (PDAC). Employing the SDSS and GAIA datasets we test the quality and ease of access to the data. PDAC will pave the way to the Science User Interface and Tools (SUIT). We employ both in-detail study of individual objects, and a statistical study of an ensemble of objects. We evaluate user-friendliness of the current interface, and make recommendations for its future improvements.

1 INTRODUCTION

This is a document to report on the user experience testing of the Preliminary Data Access Center. The Large Scale Synoptic Telescope (LSST) will produce a big volume of data. Such unprecedented data stream poses new challenges to provide an easy access for users, in such a way that they can quickly find what they need, and thus be able to focus on the science goal that they would like to achieve. The detail description of such online user-interface called Science User Interface and Tools is outlined in documents LDM-130 (SUIT requirements) and LDM-492 (SUIT Vision). An idea of having an interface to the data is not new : there exists Aladin, SDSS CAS jobs, IPAC IRSA, Mikulsky NASA Archive, NED, and many other archives. These allow a user to query for data (either via SQL query, or interface), returning the data table. Some user interfaces (eg. IRSA) have some rudimentary plotting capabilities. There have been ideas of a new interface, that would not only eg. plot the lightcurve and display the spectrum, but also allow the user to run some machine learning algorithms, or simple models that can help narrow down the query, or obtain science results in the browser. Namely, Victor Pankratius, from MIT, in his talk "Computer-Aided Discovery: Towards Scientific Insight Generation with Machine Support" outlined the idea of an ipython notebook - access to data, which lives in the cloud, is allocated some CPU share and memory, and allows one to upload / download the data and run the model in real time, which is especially helpful to geoscientists doing fieldwork, where new data acquisition conditions their next step.

Indeed, astronomers may find that quick look into the data, finding eg. all stars that exhibit RR Lyr variability and have been observed in a certain region of the sky, is very helpful.

Here we outline the user experience of PDAC (see PDAC technical description on ¹

Currently, PDAC v1, under tab 'LSST Data' in the

upper-left corner of the interface (see Fig. 1) includes the Summer 2013 DM-stack reprocessed SDSS Stripe 82 data, hosted at the NCSA on the LSST prototype ("integration cluster") hardware, in Qserv [Gregory Dubois-Felsmann, priv.comm. 02-20-2017, slack]. The reprocessing included:

- coadding the data from all epochs in each of the ugriz SDSS filters. Measurements on coadds (per object) are available as `RunDeepSource` table, accessible via Catalogs → 'DeepSource'. The single-band coadded images with MariaDB metadata are available as `DeepCoadd` table, accessible via Images → 'DeepCoadd'.
- using i-band detections to seed forced photometry on all epochs in all bands. The results of photometry are available as `RunDeepForcedSource` table, accessible via Catalogs → 'Deep Forced Source'.
- For reference, the individual calibrated single epoch images are available as `Science_Ccd_Exposure` table, accessible via Images → 'Science CCD Exposure'

Details of the S82 LSST reprocessing can be found in the PDAC document <https://confluence.lsstcorp.org/display/DM/Properties+of+the+2013+SDSS+Stripe+82+reprocessing>. Additional details of the schema are also outlined in the LSST Data Challenge Report [Shaw, Juric, Becker, Krughoff et al. 2013], and the LSST Database Schema Browser ².

PDAC v1 under tab 'External Catalogs' also provides access to all NASA/IPAC Infrared Science Archive (IRSA) publicly accessible catalogs, including GAIA, WISE, etc. (see Fig. 2). These are stored at Infrared Processing and Analysis Center (IPAC) <http://www.ipac.caltech.edu/project/lsst>.

¹ <https://confluence.lsstcorp.org/display/DM/Guide+to+PDAC+version+1>

² https://lsst-web.ncsa.illinois.edu/schema/index.php?t=DeepForcedSource&sVer=S12_lsstsim

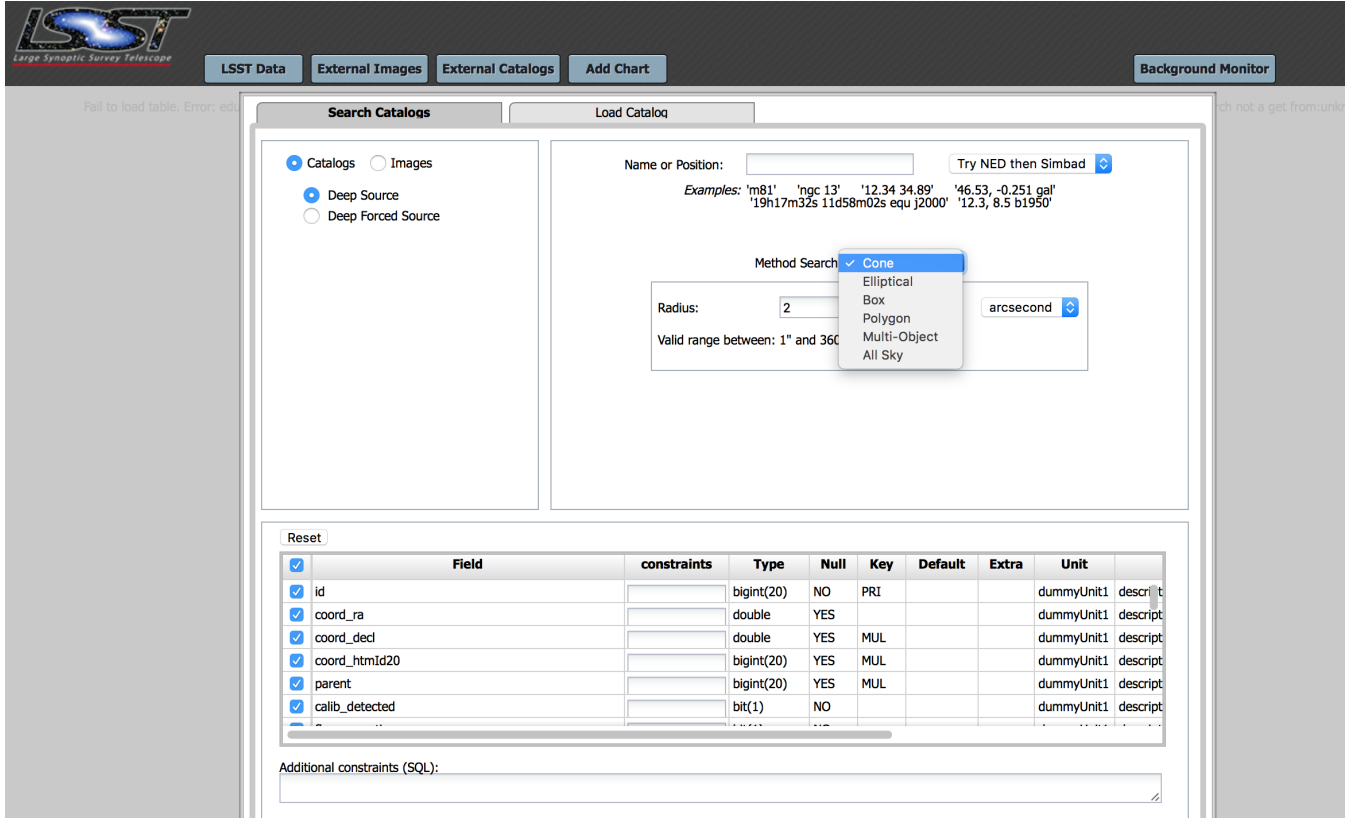


Figure 1. The main user interface of PDAC ver. 1

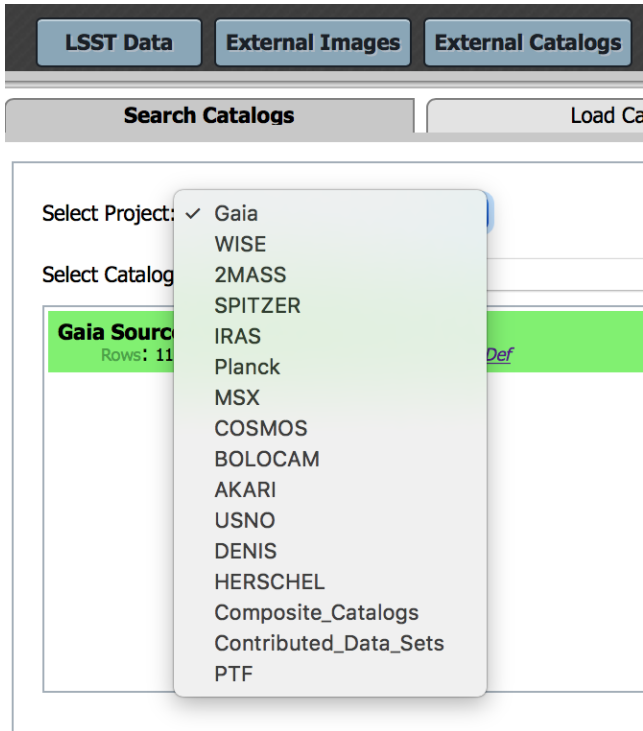


Figure 2. IPAC-hosted catalogs, accessible via IRSA.

2 METHODS

We perform single-object tests and statistical tests on an ensemble of objects.

First, we study in detail a particular source - we consider examples of variable objects, confirmed by previous studies (eg. RR Lyrae from Sesar+2010, Table 1). We download these from the S82 dataset on PDAC, run Lomb-Scargle periodogram to find period, and plot the phased lightcurve. Sesar et al. (2010) performed lightcurve template fits to 483 RR Lyrae lightcurves from SDSS (see Fig.3. Both fit parameters and lightcurves are publicly accessible in the online version of the journal

Comparing the S82 data stored at PDAC to the data from Sesar et al. (2010), we want to treat the latter as 'ground truth', but as a sanity check we perform Lomb Scargle periodogram testing to confirm the more detailed analysis of Sesar et al. (2010). Using *astroML* python module (Vanderplas et al. 2012), we sample the uniformly spaced frequency grid with $N=5000$ samples span between the smallest and the largest frequency reported in Table 1 of Sesar et al. (2010) $\pm 10\%$, i.e. $\omega_{min} = 0.9(2\pi/P_{min})$, $\omega_{max} = 1.1(2\pi/P_{min})$. We use the default *astroML* Lomb Scargle periodogram settings, namely generalized LS (see Eq.20 in Zechmeister & Kürster (2009), and Section 10.3.2 in Ivezić et al. (2014)).

Using the same frequency grid for all 483 RR Lyrae, we compute Lomb-Scargle periodograms, and determine the best-fit period from the highest frequency peak (see Fig. 5). We find that for about half of the lightcurves the Lomb-

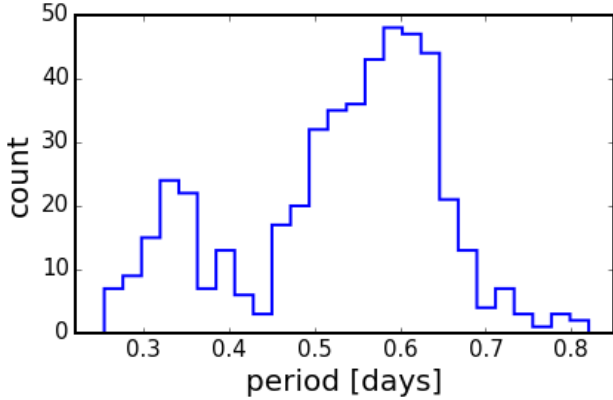


Figure 3. Distribution of RR Lyrae periods for 483 objects in (Sesar et al. 2010). Note the bimodal distribution, reflecting two main RR Lyrae types : 309 RRab (right) and 104 RRc (left) (see also Fig.16 in (Sesar et al. 2010)).

Scargle periodogram fitting single-term Fourier Series is sufficient to find the right period (note middle group centered on 1 on Fig. 5, and an example on Fig. 4). However, there are many cases where the naive single sinusoid is insufficient to correctly fit the period (groups outside of 1 on Fig. 4)

Using the Ra, dec for the RR Lyrae we positionally query the PDAC `RunDeepForcedSource` database to find objects within 2 arcsec radius. For these, we obtain calibrated g-magnitude lightcurves querying the `RunDeepForcedSource` and `Science_Ccd_Exposure` for the zero point magnitudes per exposure. Exactly as for Sesar et al. (2010) SDSS lightcurves considered before, for PDAC S82 lightcurves we also calculate Lomb-Scargle periodogram and find the most-significant frequency (fit the best period). On Fig. 9 we consider the ratio of ω_{PDAC} to ω_{true} , analogously to Fig. 5

We tested the periodogram results for few RR Lyrae using the NASA Exoplanet Archive Periodogram³. The calculation for RR Lyr ID=13350 in g-band took 15 seconds (illustrated on Fig. 10). That RR Lyr was well described by a single Fourier component - we also check an RR Lyr ID 4099 (true period 0.641754 days, SDSS LS 0.280827, PDAC LS) : period 0.64175029 , and ID 470994 (true period 0.346794 days, SDSS LS 0.531667 , PDAC LS) [MAKE A TABLE comparing SDSS-Sesar, SDSS-LS, PDAC-LS, SDSS-Exoplanet, PDAC-Exoplanet]

Second, we query the S82 database against a small subset of a given S82 patch (few degrees), downloading lightcurves for ~ 100000 objects in that area of the sky. We plot color-color diagrams, as in Sesar et al. (2007), Fig.3 ,4, and color - magnitude diagrams to show the morphology of the Sgr dSph tidal stream (Sesar et al. 2010).

3 RESULTS

4 CONCLUSIONS

ACKNOWLEDGEMENTS

Thank you !

REFERENCES

- Ivezić, Ž., Connolly, A. J., VanderPlas, J. T., & Gray, A. 2014, Statistics, Data Mining, and Machine Learning in Astronomy
 Sesar, B., et al. 2010, ApJ, 708, 717
 Sesar, B., et al. 2007, The Astronomical Journal, 134, 2236
 Vanderplas, J., Connolly, A., Ivezić, Ž., & Gray, A. 2012, in Conference on Intelligent Data Understanding (CIDU), 47
 Zechmeister, M., & Kürster, M. 2009, A&A, 496, 577

APPENDIX A: SOME EXTRA MATERIAL

If you want to present additional material which would interrupt the flow of the main paper, it can be placed in an Appendix which appears after the list of references.

This paper has been typeset from a \LaTeX file prepared by the author.

³ <http://exoplanetarchive.ipac.caltech.edu/cgi-bin/Pgram/nph-pgram>

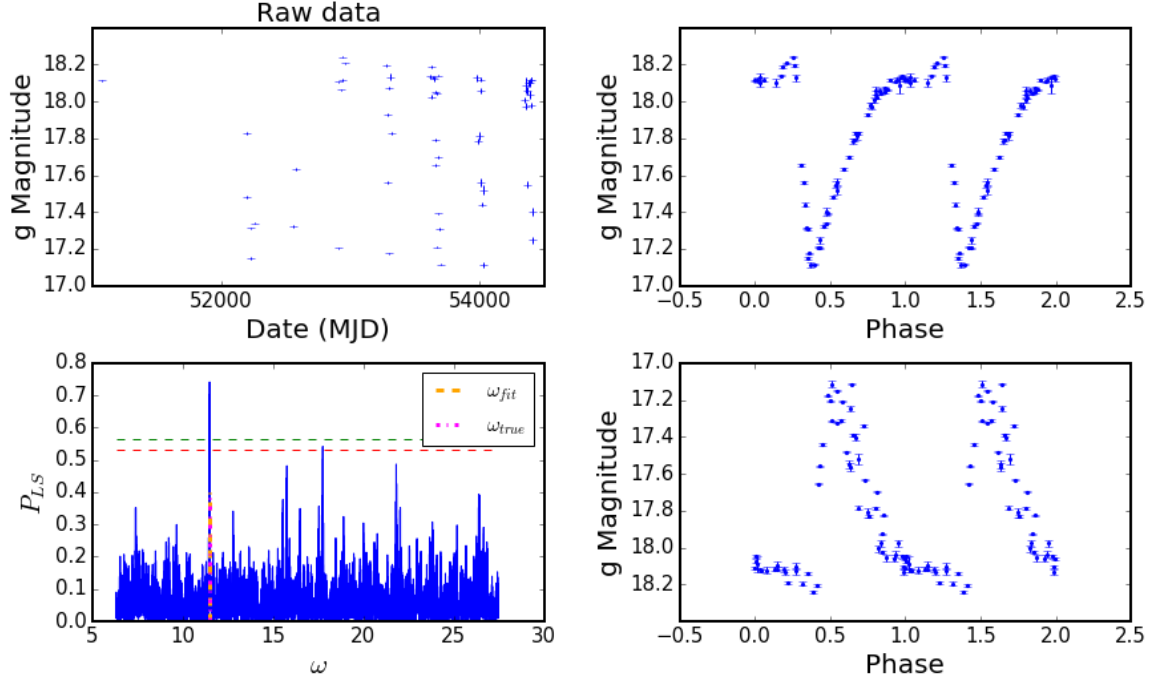


Figure 4. An example of the *astroML* Lomb Scargle periodogram performance, calculated for RR Lyr ID=13350 in SDSS g band (following Table 2 in (Sesar et al. 2010)). It took 18.6 milliseconds on a laptop to calculate this periodogram. The upper left panel depicts the raw SDSS lightcurve data. The upper right panel shows the phased lightcurve constructed with a cited period of 0.547987 days (P_{true}). The lower left panel shows the Lomb Scargle periodogram, where the orange and magenta vertical lines mark the location of the highest periodogram peak, and the frequency based on the reported period ($\omega_{true} = 2\pi/P_{true}$). The lower right panel shows the phased lightcurve constructed with the Lomb-Scargle Periodogram period of 0.547161 days, corresponding to the highest peak, $P_{fit} = 2\pi/\omega_{fit}$. The horizontal red and green lines mark the 5% and 1% significance levels for the highest peak, as found from 500 bootstrap resamplings (See http://www.astroml.org/book_figures/chapter10/index.html)

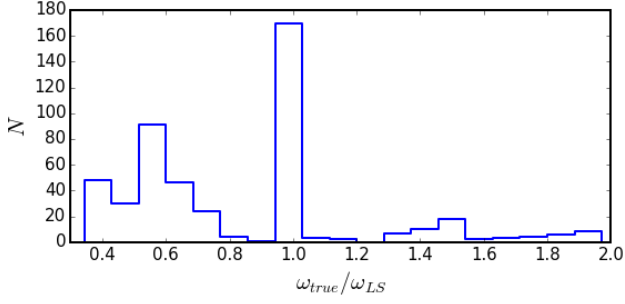


Figure 5. The distribution of the ratio of ω_{true} to ω_{fit} , where ω_{true} is inferred directly from the 'ground truth' - period cited in Table 2 of (Sesar et al. 2010). We take the same SDSS data from the paper (Table 1 in (Sesar et al. 2010)), and calculate the Lomb-Scargle single-term generalized periodogram. The frequency corresponding to the highest peak is ω_{fit} . Thus, wherever this ratio is approximately equal to 1, this means that the naive LS approach is able to recover the 'true' period. However, where the highest frequency peak is not the same as ω_{true} , the ratio will be smaller or bigger from 1. This may be caused by the inherent simplicity of the simple single-term Fourier Series fitting. Indeed, some RR Lyrae lightcurves may have shapes that are insufficiently described by a single sinusoid (as on Fig.10.18 in (Ivezić et al. 2014)).

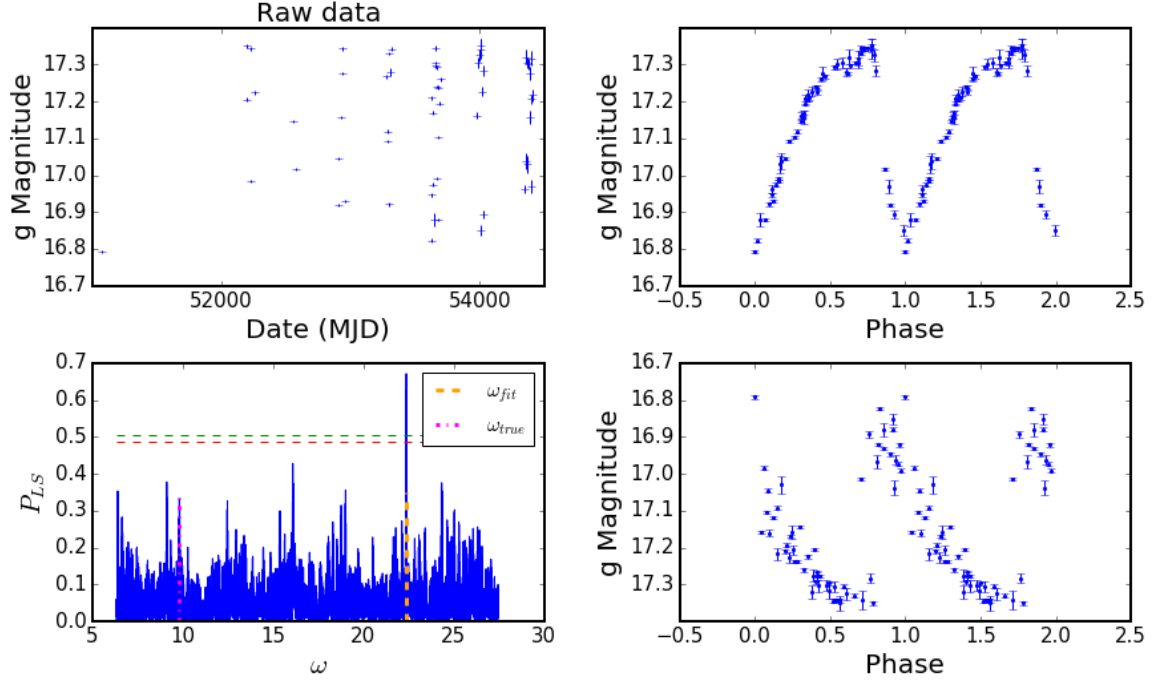


Figure 6. A failure of naive single Lomb Scargle periodogram performance - here the ratio of ω_{true} to ω_{fit} equals 0.437. Top-left : the raw SDSS lightcurve data for RR Lyr ID=4099 in g band. Top-right : the phased lightcurve constructed with a cited period of 0.641754 days (P_{true}). Bottom-left: the Lomb Scargle periodogram on a uniform frequency grid (5000 bins), with the orange and magenta vertical lines marking the location of the highest periodogram peak, and the frequency based on the reported period ($\omega_{true} = 2\pi/P_{true}$). Note that ω_{fit} and ω_{true} significantly differ for this RR Lyr, and the 'true' frequency, backed-up by the full lightcurve fitting of (Sesar et al. 2010), appears as only one of insignificant periodogram peaks. As on Fig. 4, the horizontal red and green lines mark the 5% and 1% significance levels for the highest peak, as found from 500 bootstrap resamplings. Bottom-right : the phased constructed with the P_{fit} of 0.280827 days.

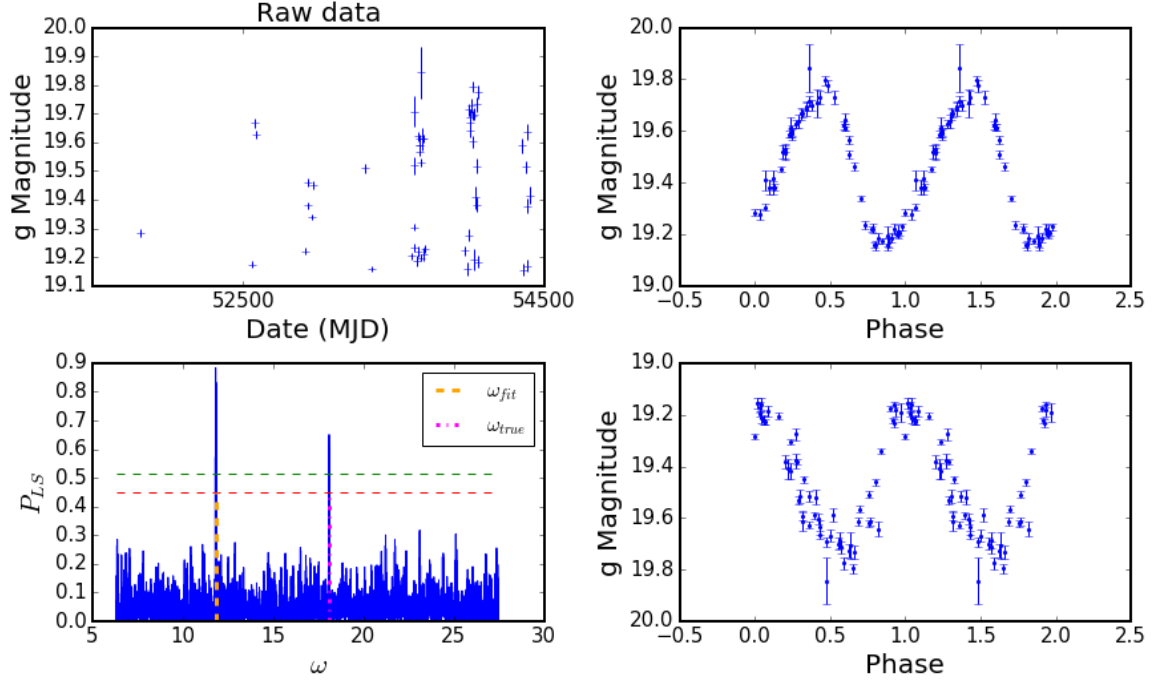


Figure 7. Same as Fig. 6, with $\omega_{true}/\omega_{fit} = 1.53$. Here RR Lyr ID=470994 has a cited period of 0.346794 days (P_{true}), whereas period derived from the Lomb-Scargle periodogram is 0.531667.

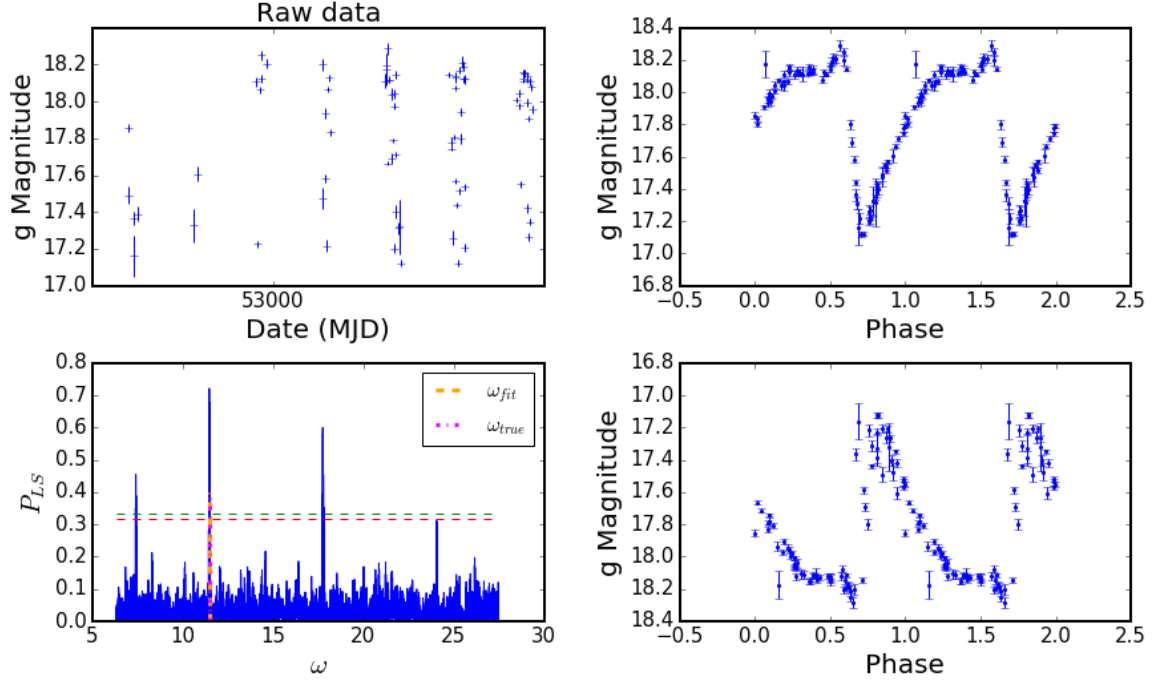


Figure 8. The same object as Fig. 4, but using data downloaded using PDAC. Using PDAC data, the RR Lyr ID=13350 has a best-fit period of 0.547969 days, almost identical to the period found by (Sesar et al. 2010) of 0.547969 days. Top-right panel shows the phased PDAC data lightcurve folded on the 'true' period, and top-right : on the 'best-fit' period. Everything else as on Fig. ??

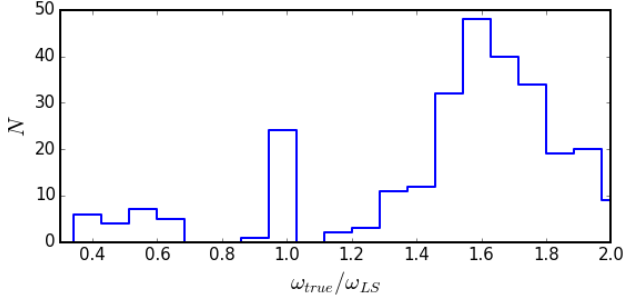


Figure 9. The distribution of the ratio of ω_{true} to ω_{PDAC} . ω_{true} is inferred directly from the 'ground truth' - period cited in the Table 2 of (Sesar et al. 2010), that used SDSS data for S82. We queried positionally against the same objects the PDAC version of SDSS reprocessed forced photometry on S82. We call this dataset 'PDAC'. For each RR Lyr star we obtained calibrated g-band lightcurves, that were used to calculate the Lomb-Scargle periodogram, on an identical frequency grid to that used previously on SDSS data from Table 1 in (Sesar et al. 2010). Based on the highest peak in each periodogram we found the best-fit frequency, ω_{PDAC} . The ratio of ω_{true} to ω_{PDAC} , analogously to ω_{true} to ω_{fit} , represents how well can naive Lomb-Scargle approach reproduce periods found by (Sesar et al. 2010) by more sophisticated lightcurve template-fitting methods. [NOTE : this is an incomplete dataset! not all RR Lyr LCs are available from PDAC → some RA, Dec have no matching PDAC object within 2 arcsecs !]

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Input File Options	Algorithm and Period Settings Reset
<p>Upload Data File: ?</p> <p>Choose File 13350_g.txt</p> <p>Upload</p> <p>Current Periodogram Data File:</p> <p>Name: 13350_g.txt</p> <p>Source: user uploaded file</p> <p>Edit Input Table</p> <p>Select Column Names:</p> <p>Time Column: col1 </p> <p>Data Column: col2 </p> <p>Plot Time vs. Data Columns</p> <p>Input File Information:</p> <p>Points used: 58 of 58</p> <p>Time range: 51075.302311 to 54412.235925</p> <p>Data range: 17.113 to 18.242</p>	<p>Select Algorithm: ?</p> <p>Algorithm: Lomb-Scargle </p> <p>Period Range:</p> <p>Minimum Period: 0.228731</p> <p>Maximum Period: 0.998246</p> <p>Period Step Method: ?</p> <p>Select Method: Fixed Frequency </p> <p>Fixed Step Size: 0.0001226</p>

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Figure 10. The same object as Fig. 4, and Fig 8, using the SDSS data from (Sesar et al. 2010). The highest significance frequency peak (power 21.58) corresponds to a period of 0.35365194 days. Only the second in significance peak (power 20.62) corresponds to the 'true' period of 0.547969 (Sesar et al. 2010). Note the bottom-left corner : the calculation took 15 secs for one lightcurve (compare to few milliseconds of AstroML code naive single-sinusoid approach that gave the same result for this particular object)/