

Crab nebula polarization observations with NIKA: implications for future CMB experiments

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Abstract

The detection of the primordial CMB B-modes in polarization constitutes one of the major challenges of modern cosmology. Their precise measurement requires accurate determination of foreground contaminants as well as of the absolute calibration of the instrument in terms of cross polarization and absolute polarization angle reconstruction. We present here a study of the Crab nebula, which is a well known polarization angle calibration source. The Crab nebula is the most intense source in the microwave sky with an extension of few arcminutes corresponding to the typical beamwidth of current CMB experiments. The Crab nebula exhibits a highly polarized synchrotron radiation at radio and millimeter wavelengths. We report in this paper high resolution (18'' FWHM) observations of the Crab nebula in intensity and polarization at 150 GHz with *NIKA*. *NIKA* operated at the IRAM 30 m telescope from 2012 to 2015, is a camera made of Lumped Element Kinetic Inductance Detectors (LEKIDs) observing the sky at 150 and 260 GHz. From these observations we are able to reconstruct the spatial distribution of the Crab nebular polarization degree and angle, which is found to be compatible with previous observations at lower and higher frequencies. Averaging across the source and using other existing data sets we find that the Crab nebula polarization angle is consistent with being constant across frequencies with a value of -87.87 ± 1.02 . We also present the first estimation of the Crab nebula Spectral Energy Distribution in polarization. These measurements will be of interest for current and next generation of CMB experiments operating at (sub)-mm and mm wavelengths.

Key words. Techniques: polarization – KIDs – individual: NIKA

1. Introduction

The polarization of the Cosmic Microwave Background (CMB) anisotropies offers a powerful way to investigate the early Universe. They can be decomposed into a scalar and a pseudo-scalar field, respectively called *E* and *B* modes. The primordial density fluctuations (scalar perturbations) can only produce *E* CMB polarization, while *B* CMB polarization can only be produced by primordial (tensor perturbations) gravitational waves (Polnarev 1985; Seljak & Zaldarriaga 1997) arising from the inflationary epoch (Guth 1981; Linde 1982) and by gravitational lensing of the *E*-modes (Planck Collaboration et al. 2016b). The detection of the primordial *B*-modes could definitively probe the existence of an inflationary epoch and constitutes one of the most ambitious goals of the modern observational cosmology.

Recently, BICEP2/Keck Collaboration et al. (2015); BICEP2 Collaboration et al. (2016) set a 95% upper limit for the detection of the tensor to scalar ratio $r < 0.07$. Upcoming CMB experiments aiming at measuring the primordial *B*-modes require an accurate determination of the foreground emissions to the CMB signal and a high control of the systematic effects. One of the most difficult parameters to be characterized for a CMB polarization experiment is the calibration of residual cross polarization and of the absolute polarization angle. This can be achieved using observations of well known polarized sources like the Crab nebula.

The Crab Nebula (or Tau A) located at equatorial coordinates $R.A. = 5^h34^m32^s$ and $Dec. = 22^\circ0'52''$ is a plerion-type supernova remnant emitting a highly polarized signal (Weiler & Panagia 1978; Michel et al. 1991). A synchrotron emission from the nebula is observed in the radio frequency domain, which is powered by a pulsar through its jet. Moreover, near the center of the nebula we observe a shock, which is formed where the jet is

thermalized and ultra-relativistic particles are released into the surrounding nebula (Weisskopf et al. 2000; Wiesemeyer et al. 2011).

This source represents the most intense polarized astrophysical object in the microwave sky at angular scales of few arcminutes and for this reason it is of particular interest for the calibration of CMB polarization experiments, which have beamwidths comparable to the extension of the source. For a more detailed review on this source we refer to Hester (2008).

The Crab nebula is used for polarization cross-check analysis in the frequency range from 30 to 353 GHz (Weiland et al. 2011; Planck Collaboration et al. 2016a). This kind of high angular resolution observations from the XPOL experiment (Thum et al. 2008) at the IRAM 30 m telescope have revealed the spatial distribution of the Crab Nebula in intensity and polarization at 90 GHz with an absolute accuracy of 0.5° in the polarization angle (Aumont et al. 2010). High angular resolution observations of the Crab nebula in polarization could be very useful for the calibration of the next generation of polarization experiments. In particular those aiming at a precise measurement of the CMB polarization, which have a large frequency range to be able to carefully study and subtract foreground emission (Kaufman et al. 2016).

Previous studies (Macías-Pérez et al. 2010) of the total flux density of the Crab nebula have shown a spectrum well described by a single synchrotron component at radio and mm wavelengths, and predict negligible variations in polarization fraction and angle in the frequency range of interest for CMB studies.

Polarized observations of the Crab Nebula have been performed with the *NIKA* camera (Monfardini et al. 2010; Catalano et al. 2014; Monfardini et al. 2014) at the IRAM 30 m tele-

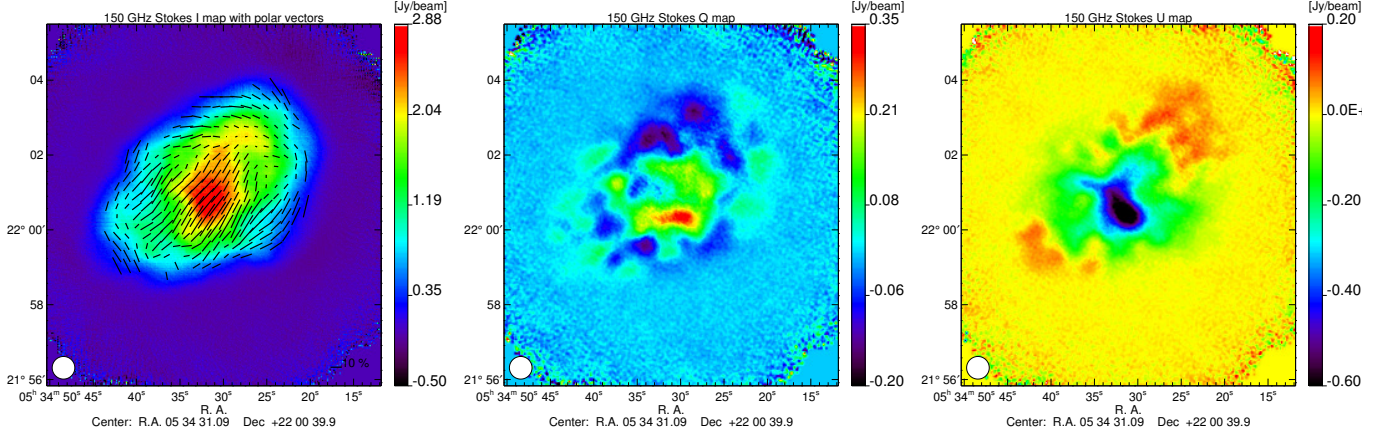


Figure 1. From left to right: Crab nebula Stokes I , Q and, U maps obtained at 150 GHz with the *NIKA* camera. Polarization vectors, indicating both the degree and the orientation, are over-plotted in black on the intensity map where the polarization intensity satisfies $P > 3 \sigma_P$, where σ_P represents the uncertainty in P .

scope during the observational campaign of February, 2015. A first overview in *NIKA* Crab polarization observations was given in Ritacco et al. (2016). In this paper we go a step forward in the analysis and combine the *NIKA* observations with previous ones in order to trace the polarized Spectral Energy Density (SED) of the Crab nebula. We use polarization observations from the WMAP satellite at 23, 33, 41, 61 and 94 GHz (Weiland et al. 2011), from the *Planck* satellite at 30, 44, 70, 100, 143, 217, 353 GHz and from XPOL at 90 GHz (Aumont et al. 2010). The paper is organized as follows: in Sec. 2 the intensity and polarization maps obtained with the *NIKA* camera are presented together with the polarization degree and angle spatial distribution; Sec. 3 presents the reconstruction of the polarization properties of the Crab nebula in well defined regions; Sec. 4 presents the Crab nebula SED in temperature and polarization; in Sec. 5 we draw conclusions.

2. NIKA observations of the Crab Nebula

2.1. NIKA camera and polarization setup

NIKA is a dual band camera observing the sky in intensity and polarization at 150 and 260 GHz with a FOV of $1.8'$, and $18''$ and $12''$ FWHM resolution, respectively. It has been operated at the IRAM 30 m telescope between 2012 and 2015. A detailed description of the *NIKA* camera can be found in Monfardini et al. (2010, 2011) and Catalano et al. (2014). In addition to intensity observations it was also a test bench for polarization facilities of the final instrument *NIKA2*, installed at the telescope in October, 2015 (Calvo et al. 2016; Catalano et al. 2016).

The polarization setup of the *NIKA* camera consists of a continuously rotating room temperature metal mesh half wave plate (HWP) followed by an analyzer at room temperature, and polarization insensitive arrays of KIDs cooled down to 100 mK. The HWP rotation modulates the input polarization signal at four times the HWP rotation frequency allowing a quasi simultaneous measurement of the I , Q and U Stokes parameters. Ritacco, A. et al. (2017) gives more details on the *NIKA* polarization facilities and describes the performance of the instrument at the telescope. *NIKA* has provided the first polarization observations performed with Kinetic Inductance Detectors, which represent

a suitable detector technology for the development of the next generation of CMB experiments.

2.2. NIKA observations

Crab nebula polarization observations with the *NIKA* camera were performed at the IRAM 30 m telescope during the observational campaign of February, 2015. Fig. 1 shows the Stokes I , Q and U maps obtained by a co-addition of 16 on-the-fly maps of 8×6 arcminutes for a total observation time of ~ 2.7 hours. The maps were performed in equatorial coordinates according to four different scan directions: 0° , 90° , 120° , 150° . This allowed us to have the best covering of the source and to reduce filtering effects.

We have used a dedicated polarization data reduction pipeline (Ritacco, A. et al. 2017), which is an extension of the intensity *NIKA* pipeline (Catalano et al. 2014; Adam et al. 2014). The main steps of the polarization pipeline are summarized below:

1. Subtraction of the HWP induced parasitic signal, which is modulated at harmonics of the HWP rotation frequency and represents the most annoying noise contributing to the polarized signal.
2. Reconstruction of the Stokes I , Q and U time ordered information (TOI) from the raw modulated data. This is achieved using a demodulation procedure consisting in a lock-in around the fourth harmonic of the HWP rotation frequency, where the polarization signal is located.
3. Subtraction of the atmospheric emission in the demodulated TOIs using decorrelation algorithms. In polarization the HWP modulation reduces significantly the atmospheric contamination and simple common mode decorrelation methods can be used. By contrast, in intensity the atmospheric emission fully dominates the signal and to recover the large angular scales we use the dual-band decorrelation algorithm presented in Adam et al. (2014).
4. Projection of the demodulated and decorrelated Stokes I , Q , and U TOIs into Stokes I , Q and U maps.
5. Correction of the intensity to polarization leakage effect, which was identified on observations of unpolarized sources like the planet Uranus. For point sources the effect is about

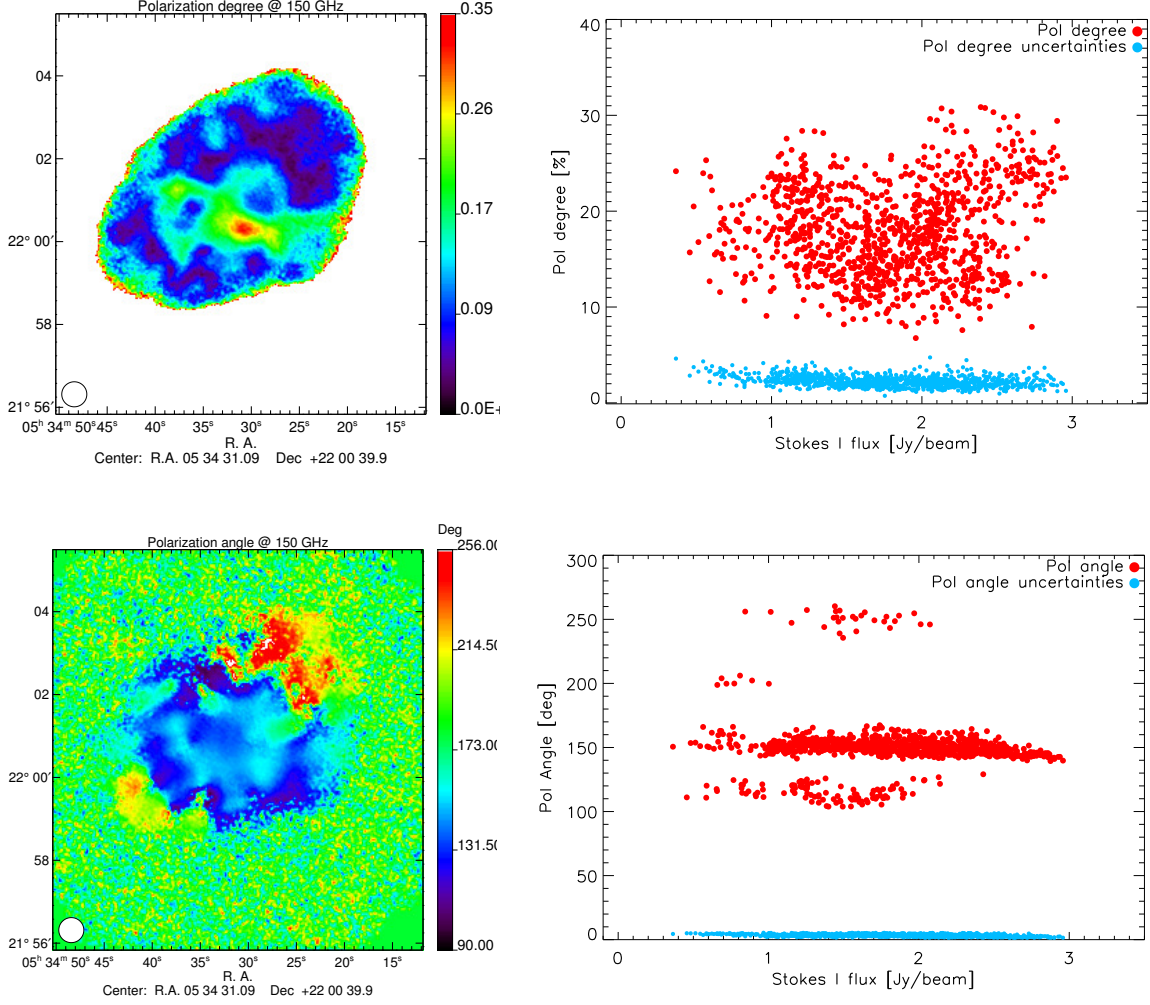


Figure 2. *Top:* noise bias uncorrected polarization degree map p (left) of the Crab nebula. The right panel shows the noise bias corrected p values as a function of total intensity map (Stokes I). The condition $P > 5\sigma_p$ is satisfied for those values. *Bottom:* polarization angle map ψ (left) of the Crab nebula. On the right panel the distribution of ψ values is represented as a function of the total intensity where $P > 5\sigma_p$. The cyan dots represent the uncertainties calculated as the dispersion between different observational scans.

3% peak-to-peak, while for extended sources like the Crab nebula is of the order of 0.5 % peak to peak. After correction the leakage effect in extended sources is negligible.

2.3. Crab polarization properties

In this section we discuss the polarization properties of the source in terms of polarization degree p and angle ψ , which are defined through the Stokes parameters I , Q , and U as follows:

$$p = \frac{\sqrt{Q^2 + U^2}}{I}$$

and

$$\psi = \frac{1}{2} \arctan \frac{U}{Q}. \quad (1)$$

These definitions are not linear in I , Q and U and are noise biased. Simmons et al. (1980); Simmons & Stewart (1985); Montier et al. (2015) proposed analytical solutions to correct this bias. At the limit of high S/N ratio the polarization degree and

its uncertainties reads:

$$p = \frac{\sqrt{Q^2 + U^2 - \sigma_Q^2 - \sigma_U^2}}{I},$$

$$\sigma_p = \frac{\sqrt{Q^2 \sigma_Q^2 + U^2 \sigma_U^2 + p^4 I^2 \sigma_I^2}}{p I^2}. \quad (2)$$

Furthermore, the polarization angle in a high S/N regime can be approximated by Eq. 1 with uncertainty

$$\sigma_\psi = \frac{\sqrt{Q^2 \sigma_Q^2 + U^2 \sigma_U^2}}{2(pI)^2}. \quad (3)$$

The polarization degree p spatial distribution of the Crab nebula without noise bias correction is presented on the top left panel of Fig. 2. The polarization reaches 35 % of the intensity across the intensity peak and decreases when moving towards the edges of the source. This feature highlights the interest of high resolution polarization observations of the Crab nebula. The

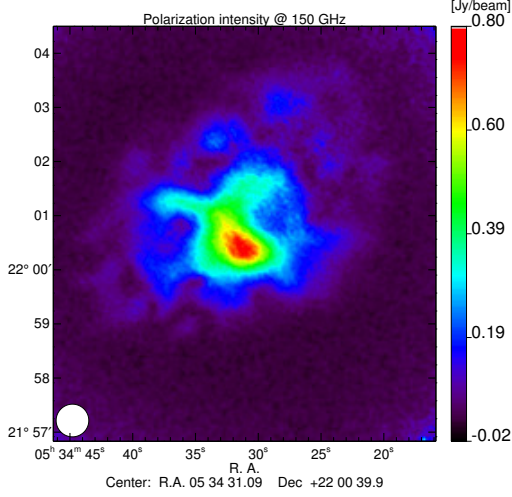


Figure 3. Polarization intensity map obtained at 150 GHz. The map shows high polarized emission reaching a value of 0.8 Jy beam^{-1} .

top right panel of the figure shows the polarization degree p as function of the total intensity map. Here the polarization degree values have been noise bias corrected and satisfy the condition $P = \sqrt{Q^2 + U^2} > 5 \sigma_P$. The distribution of the polarization degree appears highly dispersed around a mean value of 20%.

The bottom left panel of Fig. 2 shows the polarization angle, ψ , spatial distribution. We observe a relatively constant polarization angle, about 150° represented here in equatorial coordinates, except for the northern region where the averaged angle is around 220° , these values are confirmed by the bottom right panel, and some inner regions with lower polarization angle, that shows the polarization angle distribution as a function of total intensity satisfying the condition $P = \sqrt{Q^2 + U^2} > 5 \sigma_P$.

The sudden change of polarization angle on the northern region was already observed by the XPOL experiment at 90 GHz (Aumont et al. 2010). This together with the variation of the polarization fraction discussed above confirms the need of high angular resolution observations at low and high frequencies for a good understanding of the Crab polarized emission proprieties.

We present in Fig. 3 the 150 GHz Crab polarization intensity map P . We observe a peak at 0.8 Jy beam^{-1} and the polarization decreases on the extremities of the nebula.

The polarization intensity P map, represented in Fig. 3, exhibits a polarized emission peaked at 0.8 Jy beam^{-1} and decreasing to 0.1 Jy beam^{-1} on the extremities of the nebula.

3. Total intensity and polarization fluxes

We compute the total intensity flux across the Crab nebula, which has an extension of about $4'$ as shown in Fig. 1. We use standard aperture photometry techniques to calculate the flux as shown in Fig. 4. We use as center position the center of the map. A zero level in the map, calculated as the mean of the signal measured on an external annular ring region $4' < R < 4.5'$, has been subtracted from the map. The total signal estimated is $204.4 \pm 7.9 \pm 10.2 \text{ Jy}$. The first uncertainty term accounts for statistical uncertainties computed from fluctuations of the signal at large radii. The second term corresponds to calibration errors estimated to $\approx 5\%$ at 150 GHz. This absolute calibration error is

obtained from the dispersion of the estimated flux of Uranus for observations collected during the same observational campaign (Ritacco, A. et al. 2017).

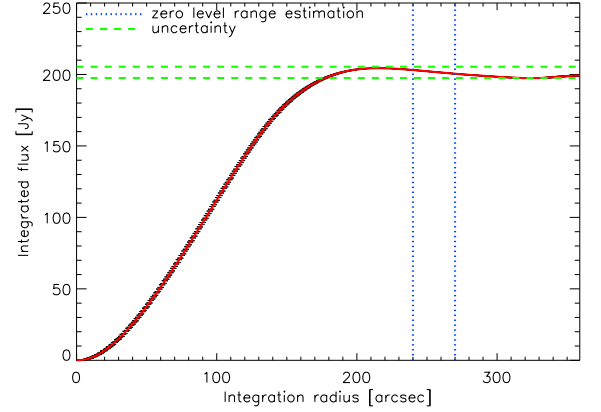


Figure 4. Cumulative flux of the Crab nebula at 150 GHz. The flux has been corrected by a zero level in the map, which corresponds to the mean of the signal calculated in an annular ring as indicated by the blue dotted lines. The green dotted line represents the uncertainties measured at large radii.

For comparison with low angular resolution CMB experiments we also give the polarization degree p and angle ψ integrated values obtained in well defined regions. Tab. 1 shows the total intensity I , polarized intensity P , polarization degree p and polarization angle ψ obtained in four different regions. To compare with the *Planck* satellite we estimate the values in two regions of $7'$ and $5'$ width corresponding to the beam widths of the *Planck* frequency channels at 143 and 217 GHz, respectively. In Tab. 1 the polarization angles are represented in Galactic coordinates to ease the comparison with the *Planck* (Planck Collaboration et al. 2016a) and *WMAP* CMB experiments (Weiland et al. 2011).

Fig. 5 shows the polarization fraction (top) and angle (bottom) of the Crab nebula as function of the frequency as measured by five different instruments: *Planck* (Planck Collaboration et al. 2016a), *WMAP* (Weiland et al. 2011), XPOL (Aumont et al. 2010), SCUPOL (Matthews et al. 2009) and NIKA (this paper). For NIKA we have averaged within a $5'$ region and considered only pixels for which $P > 3\sigma_P$. For the polarization fraction we keep the convention used by XPOL (Aumont et al. 2010) and we represent in the figure the value measured within $5'$. The solid line in both figures represents the mean value found using all the observations shown and the dashed lines represent the standard deviation.

According only for low frequency data ($< 200 \text{ GHz}$) and excluding the XPOL data we find that the degree of polarization of the Crab nebula at arcmin scales is 7.1 ± 0.1 . In terms of degree of polarization most data sets are consistent at 2σ level with the weighted-average value. For XPOL the discrepancy can probably be explained by a relatively low intensity flux with respect to the other experiences, which tends to increase the polarization fraction. In the case of *Planck* the discrepancy may be due to an evolution with frequency of the Crab nebula polarization properties. However, this hypothesis is disfavoured by polarization angle measurements which seems to be consistent with the mean value across all frequencies but for the SCUPOL data. The latter can be understood in terms of spatial variations as SCUPOL has only observed a $1.4'$ region around the peak in polarization

	I [Jy]	P [Jy]	p [%]	ψ [°]
seen by 7'	231.8±0.2	14.03±0.09	6.05±0.04	-87.14±0.07±1.8(syst)
seen by 5'	219.0±0.1	14.8 ±0.03	6.78±0.01	-87.15±0.04±1.8(syst)

Table 1. Total intensity flux I and polarized intensity flux, P , polarization degree p , and angle ψ . The values have been calculated in regions with high S/N and within 7' and 5' from the center of the source. An absolute calibration error of 5 % must be accounted for and propagated to the polarization estimates. A polarization angle systematic uncertainty of 1.8° is also considered. The polarization angle estimation satisfies the condition $P > 3\sigma_p$ to avoid noise bias.

intensity. The weighted-average of the polarization angle of the Crab nebula at arcmin scales estimated considering all available observations but SCUPOL is -88.1 ± 0.3 . All the observations shown on the bottom panel of Fig. 1 agree within 1σ with this value expect for the *Planck* value obtained at 30 GHz, which is slightly high.

4. Crab Spectral Energy Distribution in intensity and polarization

4.1. Intensity

The total intensity emission of the Crab nebula at radio and millimeter wavelengths (1 and 10^6 GHz) is mainly due to synchrotron emission and can be well described by a single power law of the form:

$$I_\nu = A(\nu/1\text{GHz})^\beta \quad (4)$$

with spectral index $\beta = -0.296 \pm 0.06$ (Baars et al. 1977; Macías-Pérez et al. 2010). At submillimeter wavelengths 10–100 μm a cold dust component is also observed (Macías-Pérez et al. 2010). Further, the Crab nebula is fading with time at a rate of $\alpha = 0.167 \pm 0.015 \text{ \% yr}^{-1}$ (Aller & Reynolds 1985). These observations suggest a low frequency emission produced by particles accelerated by the same magnetic field. The direction of the polarization is thus expected to be constant across frequencies while the polarization degree may vary.

We show in Fig. 6 the intensity flux of the Crab nebula as a function of frequency. The fluxes in the radio domain were taken from Dmitrenko et al. (1970) and Vinogradova et al. (1971). We also show microwave and mm wavelengths fluxes from *Archeops* (Macías-Pérez et al. 2007), *Planck* (Planck Collaboration et al. 2016a), WMAP (Weiland et al. 2011) and MAMBO (Bandiera et al. 2002). The measured NIKA intensity flux at 150 GHz is shown in red. We also present in the figure the best-fit to the data computed by χ^2 -minimization using observations below 100 GHz. The amplitude and spectral index found are:

$$A = 980.6 \pm 0.7, \quad \beta = -0.3151 \pm 0.0002. \quad (5)$$

For illustration we also show the best-fit to the data found by Macías-Pérez et al. (2010) in a previous analysis. Notice that fading is accounted for in the best-fit model as well as in the data represented in the figure. The estimated best-fit model (cyan line) from this paper is a slightly different w.r.t the previous analysis (green line) mainly due to the addition of recently published results by *Planck* (Planck Collaboration et al. 2016a) and WMAP (Weiland et al. 2011). However, the result found at 150 GHz by the NIKA camera agrees at 1σ level with both models. Notice that, as discussed before, XPOL intensity flux is low with respect to expectations from the two power law models.

The *Planck* data at 100, 143, 217 and 353 GHz seems to suggest another emission component in the Crab nebula. The break in the spectrum observed at around 90 GHz suggest that another population of electrons contributes to the intensity emission. By contrast the results obtained by *Archeops*, MAMBO and NIKA do not agree with this conclusion. To account for this break in the SED we assume a two power-laws model:

$$I_\nu = A(\nu/1\text{GHz})^\beta \quad \text{for } \nu < 100\text{GHz} \quad (6)$$

and

$$I_\nu = A_H(\nu/1\text{GHz})^{\beta_H} \quad \text{for } \nu > 100\text{GHz} \quad (7)$$

By χ^2 -minimization to the data we find $A_H = 8.6 \pm 0.45$ and $\beta_H = -0.71 \pm 0.09$. We note that the NIKA data is consistent with this model at the 2σ level. More data at the mm range is need to better understand the observed break of the SED at about 90 GHz.

4.2. Polarization

The total intensity of the Crab nebula has been monitored along decades across a large range of frequency by contrast the amount of polarization observations is really poor. Recent results provided by *Planck* (Planck Collaboration et al. 2016a), WMAP (Weiland et al. 2011) and XPOL (Aumont et al. 2010) together with NIKA allow us to trace the Crab nebula polarization intensity SED as shown in Fig. 7. From the figure we observe that the data are not consistent with simple power law model as for intensity. The *Planck* data show an oscillating SED while the WMAP ones are more consistent with power law model. To investigate this issue further we assume a power law model for the data as discussed by the Eq. 4.1 and fit separately the *Planck* and WMAP data. We restrict the fit to the data with $\nu < 100$ GHz. We show in Fig. 7 in black and blue the best-fit model for the *Planck* and WMAP data, respectively.

Accounting for fading the fitted amplitude and spectral indexes β are:

- best fit parameters for WMAP only data:

$$A = 78.9 \pm 7.8, \quad \beta_p = -0.35 \pm 0.03; \quad (8)$$

- best fit parameters for *Planck* only data:

$$A = 179.1 \pm 15.4, \quad \beta_p = -0.54 \pm 0.02. \quad (9)$$

The “WMAP” best-fit model is in agreement with *Planck* observations at 30, 100 and 353 GHz. The NIKA data at 150 GHz is consistent at the 1σ level with the WMAP and *Planck* best-fit models. We have also estimated the spectral index at high frequency using the map obtained by SCUPOL at 352 GHz (850 μm) and the NIKA map. Considering only the region observed by SCUPOL we obtain:

$$\beta_p = -0.33 \pm 0.01. \quad (10)$$

This result is in good agreement with the “WMAP” best-fit model.

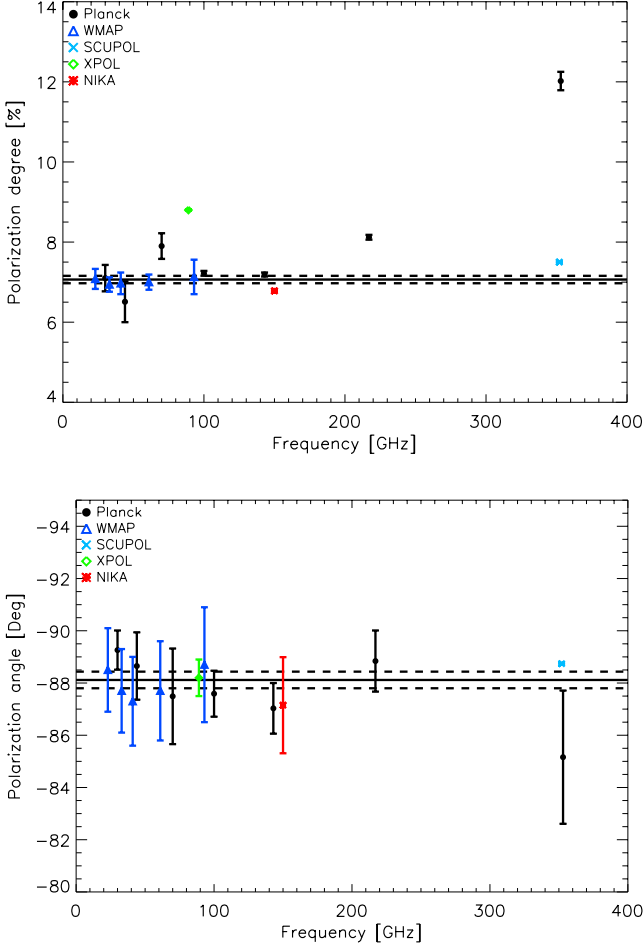


Figure 5. *Top*: polarization degree as a function of frequency as measured by *Planck* (black dots), *WMAP* (blue triangles), *XPOL* (green diamond), *SCUPOL* (cyan cross) and *NIKA* (red crosses). The *NIKA* value has been estimated by integrating in a radius of $5'$ as given by *XPOL* (Aumont et al. 2010). Solid line represents the weighted-average accounting for low frequency values (<200 GHz) and excluding *XPOL*. Dashed lines are the uncertainties. *Bottom*: polarization angles in Galactic coordinates for the same five experiments given above. Solid line represents the weighted-average accounting for all observations but *SCUPOL* and dashed lines are the uncertainties. Notice that *SCUPOL* values (Matthews et al. 2009) at 352 GHz have been estimate on maps showing only a region of $1.4'$ across the source.

5. Conclusions

The Crab nebula is considered an absolute calibrator for CMB experiments in terms of polarization degree and angle. This absolute calibration is particularly important for the measurement of the CMB polarization B-modes, which are a window towards the physics of the early Universe.

We have reported in this paper first high resolution polarization measurements of the Crab nebula at 150 GHz. These *NIKA* observations has allowed us to accurately map the spatial distribution of the polarization fraction and angle. We find an averaged polarization angle of $-87.15 \pm 0.04 \pm 1.8$ within a region of

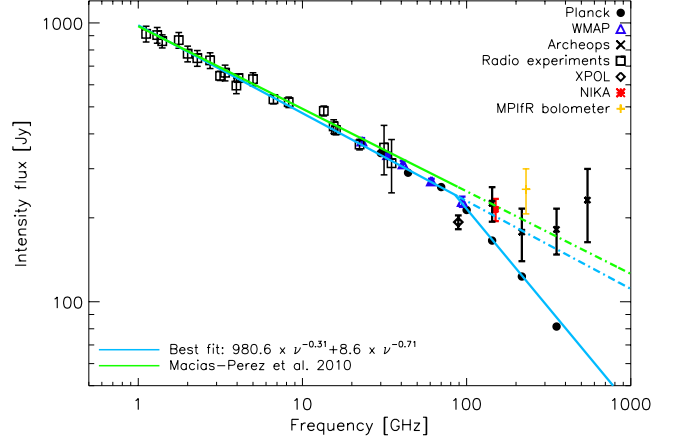


Figure 6. Crab nebula total intensity SED. The green line shows the model derived by a previous analysis discussed in (Macías-Pérez et al. 2010). The best-fit obtained by a new analysis is shown in cyan line. Both, the fit and the data account for the fading with the time.

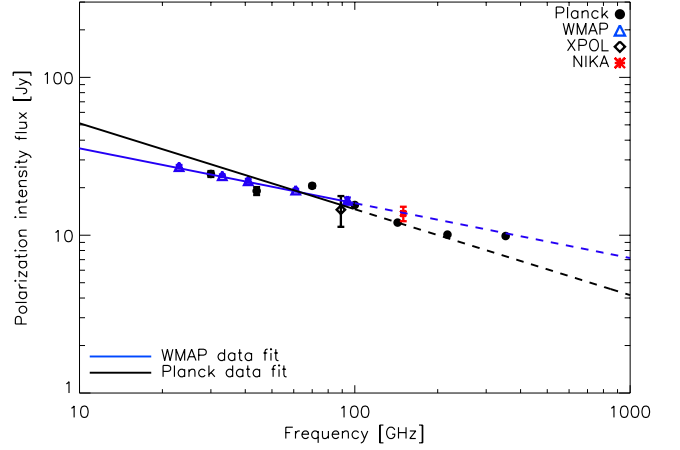


Figure 7. Crab nebula polarization intensity SED. The two best-fit models presented have been estimated using only *WMAP* data (blue line) or *Planck* data only (black line).

$5'$. This is consistent with previous measurements from *Planck*, *WMAP*, and *XPOL* with $P > 3\sigma_P$.

We conclude that the polarization angle of the Crab nebula is constant across frequencies (from 23 GHz to 217 GHz) at arcmin scales with a value of -88.1 ± 0.3 .

Using the *NIKA* data in combination with previous observations we have characterized the intensity and polarization SED of the Crab nebula. In intensity recent *Planck* data show a break in the SED suggesting the existence of two different population of electrons. In polarization we find that the data are overall consistent with a single power law spectrum as expected from synchrotron emission. However, we find some discrepancies between data sets, which will require further mm measurements at high resolution.

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