

# Crab nebula polarization observations with NIKA: implications for future CMB experiments

Address(es) of author(s) should be given

Received May 13, 2017 / Accepted –

## Abstract

The detection of the primordial CMB B-modes in polarization constitutes one of the major challenges of modern cosmology. Their precise measurement requires accurate determination of foreground contaminants as well as of the absolute calibration of the instrument in terms of cross polarization and absolute polarization angle reconstruction. We present here a study of the Crab nebula, which is a well known polarization angle calibration source. The Crab nebula is the most intense source in the microwave sky with an extension of few arcminutes corresponding to the typical beamwidth of current CMB experiments. The Crab nebula exhibits a highly polarized synchrotron radiation at radio and millimeter wavelengths. We report in this paper high resolution (18'' FWHM) observations of the Crab nebula in intensity and polarization at 150 GHz with the *NIKA* camera. *NIKA*, operated at the IRAM 30 m telescope from 2012 to 2015, is a camera made of Lumped Element Kinetic Inductance Detectors (LEKIDs) observing the sky at 150 and 260 GHz. From these observations we are able to reconstruct the spatial distribution of the Crab nebula polarization degree and angle, which is found to be compatible with previous observations at lower and higher frequencies. Averaging across the source and using other existing data sets we find that the Crab nebula polarization angle is consistent with being constant across frequencies with a value of  $-88.1 \pm 0.3$ . We also present the first estimation of the Crab nebula Spectral Energy Distribution in polarization. These measurements will be of interest for current and next generation of CMB experiments operating at microwave wavelengths.

**Key words.** Techniques: polarization – KIDs – individual: NIKA

## 1. Introduction

The polarization of the Cosmic Microwave Background (CMB) anisotropies offers a powerful way to investigate the early Universe. They can be decomposed into a scalar and a pseudo-scalar field, respectively called *E* and *B* modes. The primordial density fluctuations (scalar perturbations) can only produce *E* CMB polarization, while *B* CMB polarization can only be produced by primordial (tensor perturbations) gravitational waves (Polnarev 1985; Seljak & Zaldarriaga 1997) arising from the inflationary epoch (Guth 1981; Linde 1982) and by gravitational lensing of the *E*-modes (Planck Collaboration et al. 2016b). The detection of the primordial *B*-modes could definitively probe the existence of an inflationary epoch and constitutes one of the most ambitious goals of modern observational cosmology.

Recently, BICEP2/Keck Collaboration et al. (2015); BICEP2 Collaboration et al. (2016) set a 95% upper limit for the detection of the tensor to scalar ratio  $r < 0.07$ . Upcoming CMB experiments aiming at measuring the primordial *B*-modes require an accurate determination of the foreground emissions to the CMB signal and a high control systematic effects. One of the most difficult parameters to be characterized for a CMB polarization experiment is the calibration of residual cross polarization and of the absolute polarization angle. This can be achieved using observations of well known polarized sources like the Crab nebula.

The Crab Nebula (or Tau A) located at equatorial coordinates  $R.A. = 5^h34^m32^s$  and  $Dec. = 22^\circ0'52''$  is a plerion-type supernova remnant emitting a highly polarized signal (Weiler & Panagia 1978; Michel et al. 1991). A synchrotron emission from the nebula is observed in the radio frequency domain, which is powered by a pulsar through its jet. Moreover, near the center of the nebula we observe a shock, which is formed where the jet is

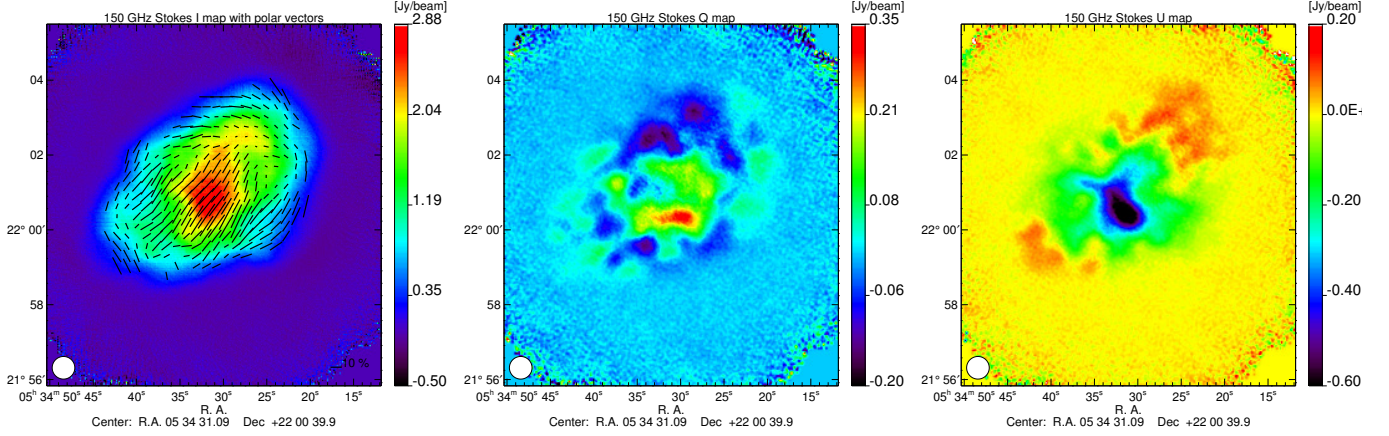
thermalized and ultra-relativistic particles are released into the surrounding nebula (Weisskopf et al. 2000; Wiesemeyer et al. 2011).

The Crab nebula is the most intense polarized astrophysical object in the microwave sky at angular scales of few arcminutes and for this reason it is of particular interest for the calibration of CMB polarization experiments, which have beamwidths comparable to the extension of the source. For a more detailed review on this source we refer to Hester (2008).

The Crab nebula is used for polarization cross-check analysis in the frequency range from 30 to 353 GHz (Weiland et al. 2011; Planck Collaboration et al. 2016a). High angular resolution observations from the XPOL experiment (Thum et al. 2008) at the IRAM 30 m telescope have revealed the spatial distribution of the Crab Nebula in intensity and polarization at 90 GHz with an absolute accuracy of  $0.5^\circ$  in the polarization angle (Aumont et al. 2010). This kind of high angular resolution observations of the Crab nebula in polarization could be very useful for the calibration of the next generation of polarization experiments. In particular those aiming at a precise measurement of the CMB polarization, which have a large frequency range to be able to carefully study and subtract foreground emission (Kaufman et al. 2016).

Previous studies (Macías-Pérez et al. 2010) of the total Spectral Energy Distribution (SED) of the Crab nebula have shown a spectrum well described by a single synchrotron component at radio and mm wavelengths, and predict negligible variations in polarization fraction and angle in the frequency range of interest for CMB studies.

Polarized observations of the Crab Nebula have been performed with the *NIKA* camera (Monfardini et al. 2010; Catalano et al. 2014; Monfardini et al. 2014) at the IRAM 30 m telescope during the observational campaign of February, 2015. A



**Figure 1.** From left to right: Crab nebula Stokes  $I$ ,  $Q$  and,  $U$  maps obtained at 150 GHz with the *NIKA* camera. Polarization vectors, indicating both the degree and the orientation, are over-plotted in black on the intensity map where the polarization intensity satisfies  $P > 3 \sigma_P$ , where  $\sigma_P$  represents the uncertainty in  $P$ .

first overview of the *NIKA* Crab polarization observations was given in Ritacco et al. (2016). In this paper we go a step forward in the analysis and combine the *NIKA* observations with previous ones in order to trace the polarized SED of the Crab nebula. We use polarization observations from the WMAP satellite at 23, 33, 41, 61 and 94 GHz (Weiland et al. 2011), from the *Planck* satellite at 30, 44, 70, 100, 143, 217, 353 GHz and from XPOL at 90 GHz (Aumont et al. 2010). The paper is organized as follows: in Sec. 2 the intensity and polarization maps obtained with the *NIKA* camera are presented together with the polarization degree and angle spatial distributions; Sec. 3 presents the reconstruction of the polarization properties of the Crab nebula in well defined regions; Sec. 4 presents the Crab nebula SED in temperature and polarization; in Sec. 5 we draw conclusions.

## 2. NIKA observations of the Crab Nebula

### 2.1. NIKA camera and polarization setup

*NIKA* is a dual band camera observing the sky in intensity and polarization at 150 and 260 GHz with a FOV of  $1.8'$ , and  $18''$  and  $12''$  FWHM resolution, respectively. It has been operated at the IRAM 30 m telescope between 2012 and 2015. A detailed description of the *NIKA* camera can be found in Monfardini et al. (2010, 2011) and Catalano et al. (2014). In addition to perform intensity observations *NIKA* was also a test bench for polarization facilities of the final instrument *NIKA2*, which was installed at the telescope in October, 2015 (Calvo et al. 2016; Catalano et al. 2016).

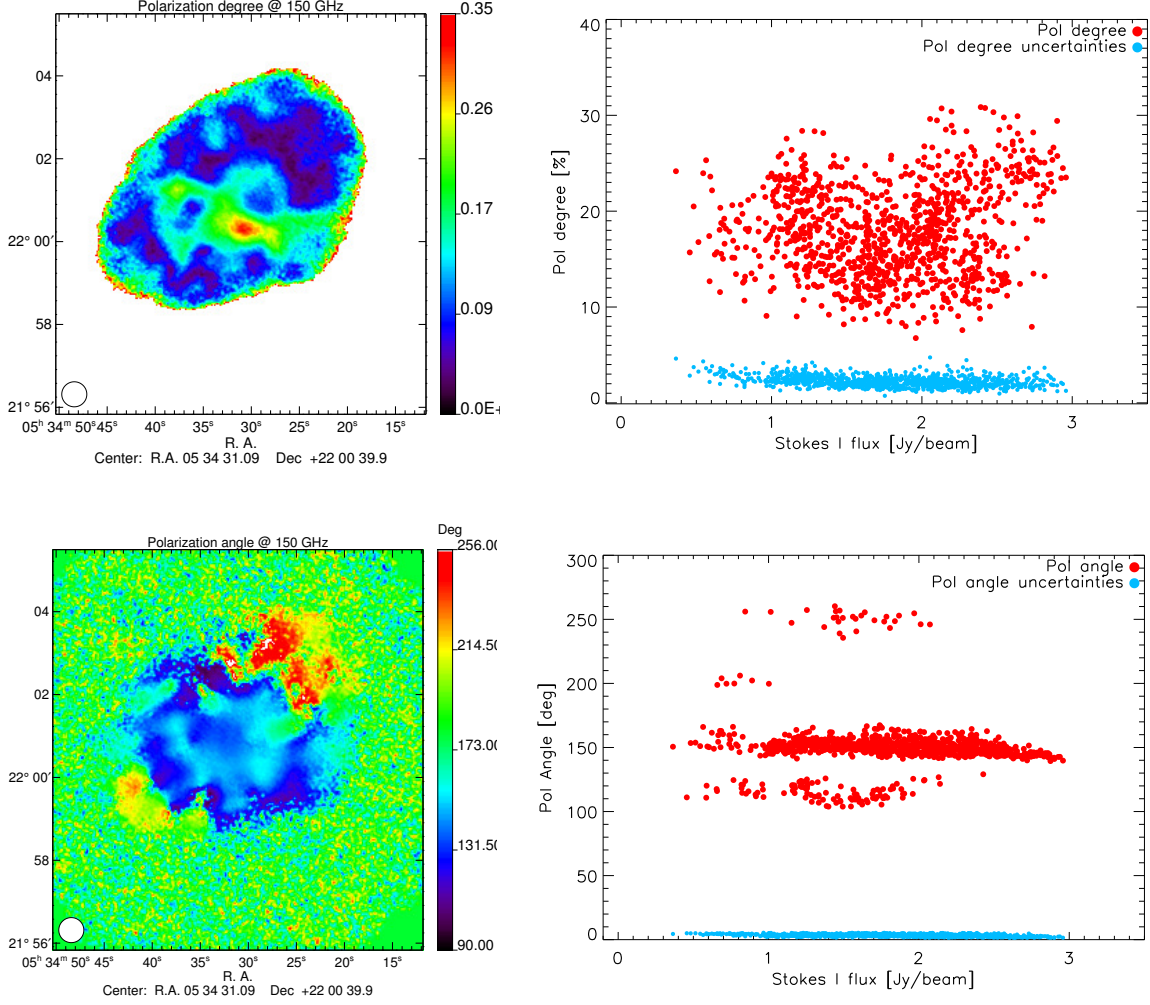
The polarization setup of the *NIKA* camera consists of a continuously rotating room temperature metal mesh half wave plate (HWP) followed by an analyzer at room temperature, and polarization insensitive arrays of KIDs cooled down to 100 mK. The HWP rotation modulates the input polarization signal at four times the HWP rotation frequency allowing a quasi simultaneous measurement of the  $I$ ,  $Q$  and  $U$  Stokes parameters. Ritacco, A. et al. (2017) gives more details on the *NIKA* polarization facilities and describes the performance of the instrument at the telescope. *NIKA* has provided the first polarization observations performed with Kinetic Inductance Detectors, which represent a suitable detector technology for the development of the next generation of CMB experiments.

### 2.2. NIKA observations

Crab nebula polarization observations with the *NIKA* camera were performed at the IRAM 30 m telescope during the observational campaign of February, 2015. Fig. 1 shows the Stokes  $I$ ,  $Q$  and  $U$  maps obtained by a co-addition of 16 on-the-fly maps of  $8 \times 6$  arcminutes for a total observation time of  $\sim 2.7$  hours. The maps were performed in equatorial coordinates according to four different scan directions:  $0^\circ$ ,  $90^\circ$ ,  $120^\circ$ ,  $150^\circ$ . This allowed us to have the best covering of the source and to reduce filtering effects.

To obtain the  $I$ ,  $Q$ , and  $U$  Crab nebula maps we have used a dedicated polarization data reduction pipeline (Ritacco, A. et al. 2017), which is an extension of the intensity *NIKA* pipeline (Catalano et al. 2014; Adam et al. 2014). The main steps of the polarization pipeline are summarized below:

1. Subtraction of the HWP induced parasitic signal, which is modulated at harmonics of the HWP rotation frequency and represents the most annoying noise contributing to the polarized signal.
2. Reconstruction of the Stokes  $I$ ,  $Q$  and  $U$  time ordered information (TOI) from the raw modulated data. This is achieved using a demodulation procedure consisting in a lock-in around the fourth harmonic of the HWP rotation frequency, where the polarization signal is located.
3. Subtraction of the atmospheric emission in the demodulated TOIs using decorrelation algorithms. In polarization the HWP modulation reduces significantly the atmospheric contamination. However simple common mode decorrelation methods can be used (Ritacco, A. et al. 2017). By contrast, in intensity the atmospheric emission fully dominates the signal and to recover the large angular scales we use the dual-band decorrelation algorithm presented in Adam et al. (2014).
4. Projection of the demodulated and decorrelated Stokes  $I$ ,  $Q$ , and  $U$  TOIs into Stokes  $I$ ,  $Q$  and  $U$  maps.
5. Correction of the intensity to polarization leakage effect, which was identified on observations of unpolarized sources like the planet Uranus. For point sources the effect is about 3% peak-to-peak, while for extended sources like the Crab nebula is of the order of 0.5 % peak to peak. After correction the leakage effect in extended sources is negligible.



**Figure 2.** *Top:* The left panel shows the noise bias uncorrected polarization degree map  $p$  of the Crab nebula. The right panel shows the noise bias corrected  $p$  values as a function of total intensity map (Stokes  $I$ ). The condition  $P > 5\sigma_P$  is satisfied for those values. *Bottom:* on the left we present the polarization angle map  $\psi$  of the Crab nebula. On the right panel the distribution of  $\psi$  values is represented as a function of the total intensity where  $P > 5\sigma_P$ . The cyan dots represent the uncertainties calculated as the dispersion between different observational scans.

### 2.3. Crab polarization properties

In this section we discuss the polarization properties of the source in terms of polarization degree  $p$  and angle  $\psi$ , which are defined through the Stokes parameters  $I$ ,  $Q$ , and  $U$  as follows:

$$p = \frac{\sqrt{Q^2 + U^2}}{I}$$

and

$$\psi = \frac{1}{2} \arctan \frac{U}{Q}. \quad (1)$$

These definitions are not linear in  $I$ ,  $Q$  and  $U$  and are noise biased. Simmons et al. (1980); Simmons & Stewart (1985); Montier et al. (2015) proposed analytical solutions to correct this bias. At the limit of high S/N ratio the polarization degree and

its uncertainties reads:

$$p = \frac{\sqrt{Q^2 + U^2 - \sigma_Q^2 - \sigma_U^2}}{I},$$

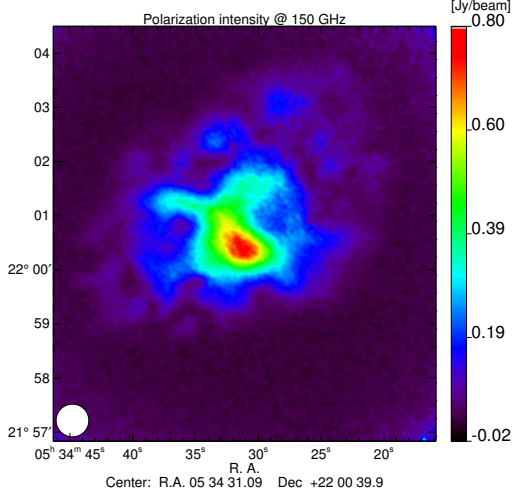
$$\sigma_p = \frac{\sqrt{Q^2 \sigma_Q^2 + U^2 \sigma_U^2 + p^4 I^2 \sigma_I^2}}{p I^2}. \quad (2)$$

Furthermore, the polarization angle in a high S/N regime can be approximated by Eq. 1 with uncertainty

$$\sigma_\psi = \frac{\sqrt{Q^2 \sigma_Q^2 + U^2 \sigma_U^2}}{2(pI)^2}. \quad (3)$$

The polarization degree  $p$  spatial distribution of the Crab nebula without noise bias correction is presented on the top left panel of Fig. 2. The polarization reaches 35 % of the intensity across the intensity peak and decreases when moving towards the edges of the source. This feature highlights the interest of high resolution polarization observations of the Crab nebula. The





**Figure 3.** NIKA Polarization intensity map of the Crab nebula at 150 GHz. The map shows high polarized emission reaching a value of  $0.8 \text{ Jy beam}^{-1}$ .

top right panel of the figure shows the polarization degree  $p$  as function of the total intensity map. Here the polarization degree values have been noise bias corrected and satisfy the condition  $P = \sqrt{Q^2 + U^2} > 5 \sigma_P$ . The distribution of the polarization degree appears highly dispersed around a mean value of 20%.

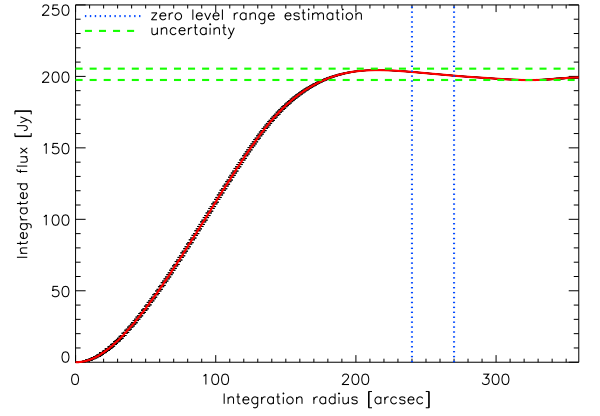
The bottom left panel of Fig. 2 shows the polarization angle,  $\psi$ , spatial distribution. We observe a relatively constant polarization angle, about  $150^\circ$  represented here in equatorial coordinates, except for the northern region where the averaged angle is around  $220^\circ$ , and some inner regions with lower polarization angle. These values are confirmed by the bottom right panel that shows the polarization angle distribution as a function of total intensity satisfying the condition  $P = \sqrt{Q^2 + U^2} > 5 \sigma_P$ .

The sudden change of polarization angle on the northern region was already observed by the XPOL experiment at 90 GHz (Aumont et al. 2010). This together with the variation of the polarization fraction discussed above confirms the need of high angular resolution observations at low and high frequencies for a good understanding of the Crab polarized emission proprieties.

We present in Fig. 3 the 150 GHz Crab polarization intensity map  $P$ . We observe a peak at  $0.8 \text{ Jy beam}^{-1}$  and the polarization decreases towards the edges of the nebula.

### 3. Total intensity and polarization fluxes

We compute the total intensity flux across the Crab nebula, which has an extension of about  $4'$  as shown in Fig. 1. We use standard aperture photometry techniques to calculate the flux as shown in Fig. 4. We use as center position the center of the map. A zero level in the map, calculated as the mean of the signal measured on an external annular ring region  $4' < R < 4.5'$ , has been subtracted from the map. The total signal estimated is  $204.4 \pm 7.9 \pm 10.2 \text{ Jy}$ . The first uncertainty term accounts for statistical uncertainties computed from fluctuations of the signal at large radii. The second term corresponds to calibration errors estimated to  $\approx 5\%$  at 150 GHz. This absolute calibration error is obtained from the dispersion of the estimated flux of Uranus for observations collected during the same observational campaign (Ritacco, A. et al. 2017).



**Figure 4.** Cumulative flux of the Crab nebula at 150 GHz. The flux has been corrected by a zero level in the map, which corresponds to the mean of the signal calculated in an annular ring as indicated by the blue dotted lines. The green dotted line represents the uncertainties measured at large radii.

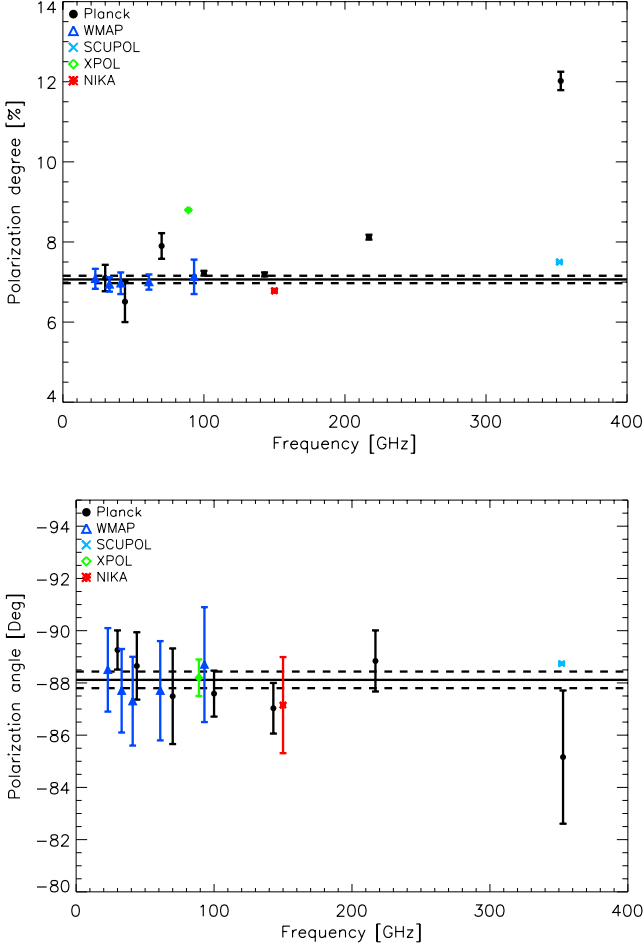
For comparison with low angular resolution CMB experiments we also give the polarization degree  $p$  and angle  $\psi$  integrated values obtained in well defined regions:  $5'$ ,  $7'$  and around the intensity peak. Tab. 1 shows the total intensity  $I$ , polarized intensity  $P$ , polarization degree  $p$  and polarization angle  $\psi$  obtained in three different regions. Notice that the polarization angles are represented in Galactic coordinates to ease the comparison with the *Planck* (Planck Collaboration et al. 2016a) and *WMAP* CMB experiments (Weiland et al. 2011).

Fig. 5 shows the polarization fraction (top) and angle (bottom) of the Crab nebula as function of the frequency as measured by five different instruments: *Planck* (Planck Collaboration et al. 2016a), *WMAP* (Weiland et al. 2011), XPOL (Aumont et al. 2010), SCUPOL (Matthews et al. 2009) and NIKA (this paper). For NIKA we have averaged within a  $5'$  region and considered only pixels for which  $P > 3\sigma_P$ . The solid line in both figures represents the weighted-average found using all the observations shown and the dashed lines represent the uncertainties estimated on the weighted-average.

Considering only low frequency data ( $< 200 \text{ GHz}$ ) and excluding the XPOL data (see below) we find that the degree of polarization of the Crab nebula at arcmin scales is  $7.1 \pm 0.1 \%$ . In terms of degree of polarization most data sets are consistent at the  $2\sigma$  level with the weighted-average value. For XPOL the discrepancy found can probably be explained by a relatively low intensity flux with respect to the other experiments, which tends to increase the polarization fraction. In the case of *Planck* the discrepancy observed at high frequency may be due to an evolution with frequency of the Crab nebula polarization proprieties. However, this hypothesis is disfavored by the polarization angle measurements that seems to be consistent with the mean value across all frequencies but for the SCUPOL data. The latter can be understood in terms of spatial variations as SCUPOL has only observed a  $1.4'$  region around the polarization intensity peak. The weighted-average of the polarization angle of the Crab nebula at arcmin scales, which was estimated considering all available observations but SCUPOL, is  $-88.1 \pm 0.3$ . All the observations shown on the bottom panel of Fig. 1 agree within  $1\sigma$  with this value except for the *Planck* value at 30 GHz, which is slightly high.

|  | $I$<br>[Jy] | $P$<br>[Jy] | $p$<br>[%] | $\psi$<br>[°]         |
|--|-------------|-------------|------------|-----------------------|
| 7' centered  | 231.8±0.2   | 14.03±0.09  | 6.05±0.04  | -87.14±0.07±1.8(syst) |
| 5' centered  | 219.0±0.1   | 14.8±0.03   | 6.78±0.01  | -87.15±0.04±1.8(syst) |
| 1.4' polarization intensity peak with $P > 0.2$ Jy | 60.6±0.2    | 10.28±0.01  | 16.96±0.01 | -87.69±0.01±1.8(syst) |

**Table 1.** Total intensity flux  $I$ , polarized intensity flux  $P$ , polarization degree  $p$ , and angle  $\psi$ . The values have been calculated in regions with high S/N and within 7', 5' radius from the center of the source, and around the peak of intensity in a region of 1.4' radius. An absolute calibration error of 5 % must be accounted for and propagated to the polarization estimates. A polarization angle systematic uncertainty of 1.8° is also considered. The polarization angle estimation satisfies the condition  $P > 3\sigma_P$  to avoid noise bias.



**Figure 5.** *Top:* polarization degree as a function of frequency as measured by *Planck* (black dots), *WMAP* (blue triangles), *XPOL* (green diamond), *SCUPOL* (cyan cross) and *NIKA* (red crosses). The *NIKA* value has been estimated by integrating in a radius of 5' as given by *XPOL* (Aumont et al. 2010). The solid line represents the weighted-averaged degree of polarization accounting for low frequency values (<200 GHz) and excluding *XPOL*. Dashed lines are the uncertainties. *Bottom:* polarization angles in Galactic coordinates for the same five experiments given above. The solid line represents the weighted-averaged polarization angle accounting for all observations but *SCUPOL* and the dashed lines are the uncertainties. Notice that *SCUPOL* values (Matthews et al. 2009) at 352 GHz have been estimated on maps covering only a region of 1.4' radius centered in the peak polarization intensity.

## 4. Crab Spectral Energy Distribution in intensity and polarization

### 4.1. Intensity

The total intensity emission of the Crab nebula at radio and millimeter wavelengths (1 and 500 GHz) is mainly due to synchrotron emission and can be well described by a single power law of the form:

$$I_\nu = A(\nu/1\text{GHz})^\beta \quad (4)$$

with spectral index  $\beta = -0.296 \pm 0.06$  (Baars et al. 1977; Macías-Pérez et al. 2010). Further, the Crab nebula is fading with time at a rate of  $\alpha = 0.167 \pm 0.015 \text{ \% yr}^{-1}$  (Aller & Reynolds 1985). These results suggest a low frequency emission produced by particles accelerated by the same magnetic field. The direction of the polarization is therefore expected to be constant across frequencies while the polarization degree may vary.

We show in Fig. 6 the intensity flux of the Crab nebula as a function of frequency. The fluxes in the radio domain were taken from Dmitrenko et al. (1970) and Vinogradova et al. (1971). We also show microwave and mm wavelengths fluxes from *Archeops* (Macías-Pérez et al. 2007), *Planck* (Planck Collaboration et al. 2016a), *WMAP* (Weiland et al. 2011) and *MAMBO* (Bandiera et al. 2002). The measured *NIKA* intensity flux at 150 GHz is shown in red. We also present in the figure the best-fit to the data computed by  $\chi^2$ -minimization using observations below 100 GHz (cyan line). The amplitude and spectral index found are:

$$A = 980.6 \pm 0.7, \quad \beta = -0.3151 \pm 0.0002. \quad (5)$$

For illustration we also show the best-fit model to the data found by Macías-Pérez et al. (2010) in a previous analysis (green line). Notice that fading is accounted for in the best-fit models as well as in the data presented in the figure. The estimated best-fit model from this paper is slightly different w.r.t the previous analysis mainly due to the addition of recently published results by *Planck* (Planck Collaboration et al. 2016a) and *WMAP* (Weiland et al. 2011). However, the result found at 150 GHz by the *NIKA* camera agrees at  $1\sigma$  level with both models. Notice that, as discussed before, the *XPOL* intensity flux is low with respect to expectations from the two power law models.

The *Planck* data at 100, 143, 217 and 353 GHz are consistent with a much steeper SED with respect the lower frequency data suggesting an extra emission component in the Crab nebula. This emission could be due to a different population of relativistic electrons (?). Notice that a break in the SED of the Crab Nebula is also found for frequencies above 3000 GHz (Macías-Pérez et al. 2010). By contrast, the results obtained by *Archeops*, *MAMBO* and *NIKA* do not seem to support such a low frequency

break in the Crab nebula SED. However, to account for this break in the SED we assume a simple two-power-laws model<sup>1</sup>:

$$I_\nu = A(\nu/1\text{GHz})^\beta \quad \text{for } \nu < 100\text{GHz} \quad (6)$$

and

$$I_\nu = A_H(\nu/1\text{GHz})^{\beta_H} \quad \text{for } \nu > 100\text{GHz} \quad (7)$$

By  $\chi^2$ -minimization we find  $A_H = 8.6 \pm 0.45$  and  $\beta_H = -0.71 \pm 0.09$ . We note that the *NIKA* data is consistent with this model at the  $2\sigma$  level. More data at mm wavelengths are needed to better understand the observed break of the SED at about 90 GHz.

#### 4.2. Polarization

The total intensity of the Crab nebula has been monitored along decades across a large range of frequencies. By contrast the amount of polarization observations is really poor. Recent results provided by *Planck* (Planck Collaboration et al. 2016a), *WMAP* (Weiland et al. 2011) and *XPOL* (Aumont et al. 2010) together with *NIKA* allow us to trace the Crab nebula polarization intensity SED as shown in Fig. 7. From the figure we observe that the data are not consistent with a single power law model as for intensity. The *Planck* data points show an oscillating behavior while the *WMAP* ones are more consistent with a power law model. Furthermore, the average slope of the SED when approached by a single power-law seems to be different for the *Planck* and *WMAP* data. To investigate this issue further we assume a power law model for the data as in Eq. 4.1, and fit separately the *Planck* and *WMAP* data. We restrict the fit to the data with  $\nu < 100$  GHz. We show in Fig. 7 in black and blue the best-fit models for the *Planck* and *WMAP* data, respectively.

Accounting for fading the fitted amplitudes and spectral indexes are:

- best-fit parameters for *WMAP* only data:

$$A = 78.9 \pm 7.8 \quad , \quad \beta_p = -0.35 \pm 0.03; \quad (8)$$

- best-fit parameters for *Planck* only data:

$$A = 179.1 \pm 15.4 \quad , \quad \beta_p = -0.54 \pm 0.02. \quad (9)$$

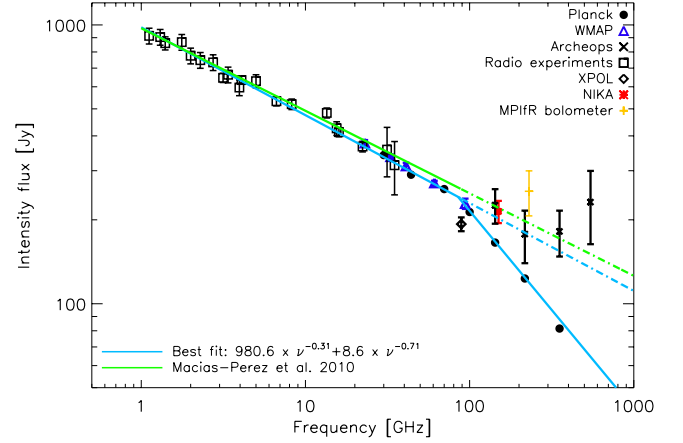
The *WMAP* best-fit model is in agreement with *Planck* observations at 30, 100 and 353 GHz. The *Planck* best-fit model seems to overestimate the polarization flux at low frequency. The *NIKA* data at 150 GHz are consistent at the  $1\sigma$  level with both the *WMAP* and *Planck* best-fit models. In contrast to the intensity data, we find no clear hint of a break in the polarisation SED at frequencies above 90 GHz.

We have also estimated the spectral index of the Crab nebula polarization emission at high frequency using the map obtained by SCUPOL at 352 GHz (850  $\mu\text{m}$ ) and the *NIKA* map. Considering only the region observed by SCUPOL we obtain:

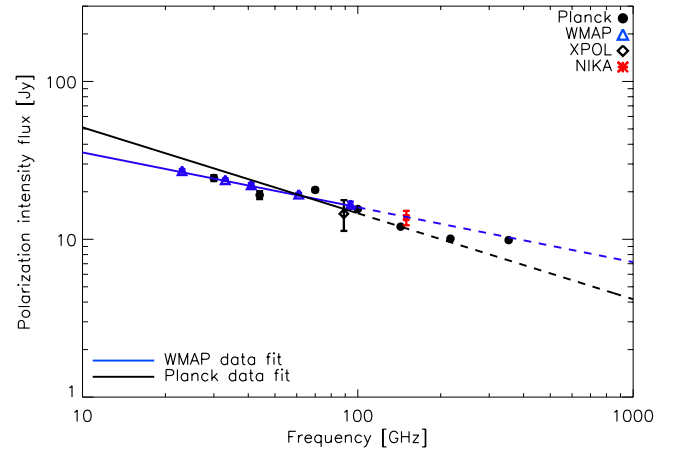
$$\beta_p = -0.33 \pm 0.01. \quad (10)$$

This result is in good agreement with the *WMAP* best-fit model spectral index.

<sup>1</sup> A detailed physical description of the SED of the Crab nebula is out of the scope of this paper.



**Figure 6.** Crab nebula total intensity SED as obtained from *Planck* (Planck Collaboration et al. 2016a), *WMAP* (Weiland et al. 2011), *Archeops* (Macías-Pérez et al. 2007), radio experiments (Dmitrenko et al. 1970; Vinogradova et al. 1971), *XPOL* (Aumont et al. 2010), *NIKA* (this paper) and *MAMBO* (Bandiera et al. 2002) data. The green line shows the model derived by a previous analysis discussed in (Macías-Pérez et al. 2010). The best-fit obtained by the analysis in this paper is shown in cyan line. Both, the best-fit models and the data account for the Crab nebula fading with the time.



**Figure 7.** Crab nebula polarization intensity SED as obtained from the *Planck* (Planck Collaboration et al. 2016a), *WMAP* (Weiland et al. 2011) and *NIKA* (this paper) data. The two best-fit models presented have been estimated using only *WMAP* data (blue line) or *Planck* data only (black line).

## 5. Conclusions

The Crab nebula is considered as an absolute calibrator for CMB experiments in terms of polarization degree and angle. This absolute calibration is particularly important for the measurement of the CMB polarization B-modes, which are a window towards the physics of the early Universe.

We have reported in this paper first high resolution polarization measurements of the Crab nebula at 150 GHz with the *NIKA* camera. These observations have allowed us to accurately map

the spatial distribution of the polarization fraction and angle. We find an averaged polarization angle of  $-87.15 \pm 0.04 \pm 1.8$  within a region of  $5'$  radius with respect to the central coordinates of the map R.A.:  $05^h 34^m 31.09$  s, Dec.:  $+22^\circ 00' 39.9''$ . This is consistent with previous measurements from *Planck*, *WMAP*, and *XPOL*.

Using all available data sets in polarization to date we conclude that the polarization angle of the Crab nebula is consistent with being constant across frequencies, from 23 GHz to 217 GHz, at arcmin scales with a value of  $-88.1 \pm 0.3$ .

Moreover we have characterized the intensity and polarization SED of the Crab nebula. In intensity recent *Planck* data show a break in the SED suggesting the existence of two different populations of relativistic electrons. In polarization we find that the data are overall consistent with a single power law spectrum as expected from synchrotron emission. However, we find some discrepancies between data sets, which will require further mm measurements at high resolution for a better understanding of the physics at play.

**Acknowledgements.** We would like to thank the IRAM staff for their support during the *NIKA* campaigns. The NIKA dilution cryostat has been designed and built at the Institut Néel. In particular, we acknowledge the crucial contribution of the Cryogenics Group, and in particular Gregory Garde, Henri Rodenas, Jean Paul Leggeri, Philippe Camus. This work has been partially funded by the Foundation Nanoscience Grenoble, the LabEx FOCUS ANR-11-LABX-0013 and the ANR under the contracts “MKIDS”, “NIKA” and ANR-15-CE31-0017. This work has benefited from the support of the European Research Council Advanced Grant ORISTARS under the European Union’s Seventh Framework Programme (Grant Agreement no. 291294). We acknowledge fundings from the ENIGMASS French LabEx (R. A. and F. R.), the CNES post-doctoral fellowship program (R. A.), the CNES doctoral fellowship program (A. R.) and the FOCUS French LabEx doctoral fellowship program (A. R.).

## References

- Adam, R., Comis, B., Macías-Pérez, J. F., et al. 2014, *A&A*, 569, A66  
 Aller, H. & Reynolds, S. 1985, *The Astrophysical Journal*, 293, L73  
 Aumont, J., Conversi, L., Thum, C., et al. 2010, *A&A*, 514, A70  
 Baars, J., Genzel, R., Pauliny-Toth, I., & Witzel, A. 1977, *Astronomy and Astrophysics*, 61, 99  
 Bandiera, R., Neri, R., & Cesaroni, R. 2002, *A&A*, 386, 1044  
 BICEP2 Collaboration, Keck Array Collaboration, Ade, P. A. R., et al. 2016, *Physical Review Letters*, 116, 031302  
 BICEP2/Keck Collaboration, Planck Collaboration, Ade, P. A. R., et al. 2015, *Physical Review Letters*, 114, 101301  
 Calvo, M., Benoît, A., Catalano, A., et al. 2016, *Journal of Low Temperature Physics*, 184, 816  
 Catalano, A., Adam, R., Ade, P., et al. 2016, *ArXiv e-prints*  
 Catalano, A., Calvo, M., Ponthieu, N., et al. 2014, *A&A*, 569, A9  
 Dmitrenko, D., Tseitlin, N., Vinogradova, L., & Giterman, K. F. 1970, *Radiophysics and Quantum Electronics*, 13, 649  
 Guth, A. H. 1981, *Phys. Rev. D*, 23, 347  
 Hester, J. J. 2008, *ARA&A*, 46, 127  
 Kaufman, J., Leon, D., & Keating, B. 2016, *International Journal of Modern Physics D*, 25, 1640008  
 Linde, A. D. 1982, *Physics Letters B*, 108, 389  
 Macías-Pérez, J., Lagache, G., Maffei, B., et al. 2007, *Astronomy & Astrophysics*, 467, 1313  
 Macías-Pérez, J. F., Mayet, F., Aumont, J., & Désert, F.-X. 2010, *ApJ*, 711, 417  
 Matthews, B. C., McPhee, C. A., Fissel, L. M., & Curran, R. L. 2009, *ApJS*, 182, 143  
 Michel, F. C., Scowen, P. A., Dufour, R. J., & Hester, J. J. 1991, *ApJ*, 368, 463  
 Monfardini, A., Adam, R., Adane, A., et al. 2014, *Journal of Low Temperature Physics*, 176, 787  
 Monfardini, A., Benoit, A., Bideaud, A., et al. 2011, *ApJS*, 194, 24  
 Monfardini, A., Swenson, L. J., Bideaud, A., et al. 2010, *A&A*, 521, A29  
 Montier, L., Plaszczyński, S., Levrier, F., et al. 2015, *A&A*, 574, A136  
 Planck Collaboration, Ade, P. A. R., Aghanim, N., et al. 2016a, *A&A*, 594, A26  
 Planck Collaboration, Ade, P. A. R., Aghanim, N., et al. 2016b, *A&A*, 596, A102  
 Polnarev, A. 1985, *Soviet Astronomy*, 29, 607  
 Ritacco, A., Adam, R., Adane, A., et al. 2016, *Journal of Low Temperature Physics*, 184, 724  
 Ritacco, A., Ponthieu, N., Catalano, A., et al. 2017, *A&A*, 599, A34  
 Seljak, U. & Zaldarriaga, M. 1997, *Physical Review Letters*, 78, 2054  
 Simmons, J. F. L., Aspin, C., & Brown, J. C. 1980, *A&A*, 91, 97  
 Simmons, J. F. L. & Stewart, B. G. 1985, *A&A*, 142, 100  
 Thum, C., Wiesemeyer, H., Paubert, G., Navarro, S., & Morris, D. 2008, *PASP*, 120, 777  
 Vinogradova, L. V., Dmitrenko, D. A., & Tseitlin, N. M. 1971, *Izvestiia Vysshiaia Uchebn. Zaved., Radiofizika*, 14, 157  
 Weiland, J. L., Odegard, N., Hill, R. S., et al. 2011, *ApJS*, 192, 19  
 Weiler, K. W. & Panagia, N. 1978, *A&A*, 70, 419  
 Weisskopf, M. C., Hester, J. J., Tennant, A. F., et al. 2000, *ApJ*, 536, L81  
 Wiesemeyer, H., Thum, C., Morris, D., Aumont, J., & Rosset, C. 2011, *A&A*, 528, A11