

# Crab nebula polarization observations at 150 GHz with NIKA: implications for future CMB experiments

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## Abstract

The detection of the primordial Cosmic Microwave Background (CMB) polarization B-modes constitutes one of the major challenges of modern cosmology. Their precise measurement requires an accurate determination of foreground contaminants as well as of the absolute calibration of the instrument in terms of cross polarization and absolute polarization angle reconstruction. We present here a study of the Crab nebula, a supernova remnant, which is a well known calibration source in temperature and commonly used in polarization. The Crab nebula is the brightest source in the microwave sky with an extension of a few arcminutes corresponding to the typical beam-width of current CMB experiments. The Crab nebula exhibits highly polarized synchrotron radiation at radio and millimeter wavelengths. In addition to CMB measurements, it can be used as standard calibrator for any experiment which aims at measuring the polarization of the sky. We report in this paper high resolution (18'' FWHM) observations of the Crab nebula in total intensity and linear polarization at 150 GHz with the *NIKA* camera. *NIKA*, operated at the IRAM 30 m telescope from 2012 to 2015, is a camera made of Lumped Element Kinetic Inductance Detectors (LEKIDs) observing the sky at 150 and 260 GHz. From these observations we are able to reconstruct the spatial distribution of the Crab nebula polarization degree and angle, which is found to be compatible with previous observations at lower and higher frequencies. Averaging across the source and using other existing data sets we find that the Crab nebula polarization angle is consistent with being constant over a wide range of frequencies with a value of  $-88.1^\circ \pm 0.3$  in Galactic coordinates. We also present the first estimation of the Crab nebula Spectral Energy Distribution polarized flux. These measurements will be crucial to calibrate current and next generation of CMB experiments operating at microwave wavelengths.

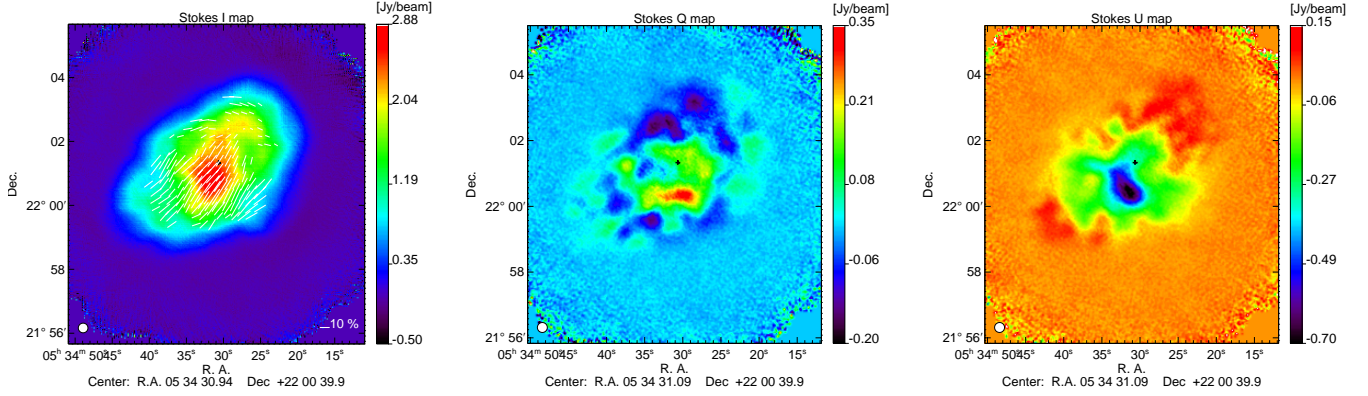
**Key words.** Techniques: Crab nebula – Tau A – polarization – KIDs – individual: NIKA

## 1. Introduction

The polarization of the Cosmic Microwave Background (CMB) anisotropies offers a powerful way to investigate the early Universe. It can be decomposed into a scalar and a pseudo-scalar

field, respectively called *E* and *B* modes. The primordial density fluctuations (scalar perturbations) can only produce *E* type CMB polarization, while *B* type CMB polarization can only be produced by primordial (tensor perturbations) gravitational waves (Polnarev 1985; Seljak & Zaldarriaga 1997) arising from the inflationary epoch (Guth 1981; Linde 1982) and by gravitational lensing of the *E*-modes (Planck Collaboration et al. (2016b) and

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**Figure 1.** From left to right: Crab nebula Stokes  $I$ ,  $Q$  and,  $U$  maps shown here in Equatorial coordinates obtained at 150 GHz with the *NIKA* camera. Polarization vectors, indicating both the degree and the orientation, are over-plotted in white on the intensity map where the polarization intensity satisfies  $I_{pol} > 3\sigma_{I_{pol}}$  and  $I_{pol} > 0.1$  Jy. The *NIKA* FWHM is shown in the lower left. The black cross marks the pulsar position.

refs. therein). The detection of the primordial  $B$ -modes **would set tight constraints on Inflation models** and constitutes one of the most ambitious goals of modern observational cosmology.

Recently, [BICEP2/Keck Collaboration et al. \(2015\)](#); [BICEP2 Collaboration et al. \(2016\)](#) set a 95% upper limit for the detection of the tensor to scalar ratio  $r < 0.07$ . Upcoming CMB experiments aiming at measuring the primordial  $B$ -modes require an accurate determination of the foreground emissions to the CMB signal and a high control of systematic effects. One of the most difficult parameters to be characterized for a CMB polarization experiment is the calibration of residual cross polarization and of the absolute polarization angle. This can be achieved using observations of well known polarized sources like the Crab nebula. The Crab Nebula (or Tau A) is a plerion-type supernova remnant emitting a highly polarized signal ([Weiler & Panagia 1978](#); [Michel et al. 1991](#)). The synchrotron emission from the nebula is observed in the radio frequency domain, which is powered by a pulsar located at equatorial coordinates (J2000)  $R.A. = 5^h34^m32^s$  and  $Dec. = 22^\circ0'52''$  through its jet. Moreover, near the center of the nebula one observes a shock, which is formed where the jet is thermalized and ultra-relativistic particles are released into the surrounding nebula ([Weisskopf et al. 2000](#); [Wiesemeyer et al. 2011](#)). The Crab nebula is the most intense polarized astrophysical object in the microwave sky at angular scales of few arcminutes and for this reason it is of particular interest for the calibration of CMB polarization experiments, which have beamwidths comparable to the extension of the source. It is also quite isolated with low background diffuse emission. For a more detailed review on this source we refer to [Hester \(2008\)](#).

The Crab nebula is used for polarization cross-check analysis in the frequency range from 30 to 353 GHz ([Weiland et al. 2011](#); [Planck Collaboration et al. 2016a](#)). High angular resolution observations from the XPOL experiment ([Thum et al. 2008](#)) at the IRAM 30 m telescope have revealed the spatial distribution of the Crab Nebula in intensity and polarization at 90 GHz with an absolute accuracy of  $0.5^\circ$  in the polarization angle ([Aumont et al. 2010](#)). Such high resolution observations of the Crab nebula in polarization are crucial for future *NIKA2* polarization observations and could be very useful for the calibration of the next generation of polarization experiments, in particular those aiming at a precise measurement of the CMB polarization that have

a large frequency range to be able to carefully study and subtract foreground emission ([Kaufman et al. 2016](#)).

Previous studies ([Macías-Pérez et al. 2010](#)) of the total Spectral Energy Distribution (SED) of the Crab nebula have shown a spectrum well described by a single synchrotron component at radio and mm wavelengths and predict negligible variations in polarization fraction and angle in the frequency range of interest for CMB studies.

Observations of the polarization of the Crab Nebula have been performed with the *NIKA* camera ([Monfardini et al. 2010](#); [Catalano et al. 2014](#); [Monfardini et al. 2014](#)) at the IRAM 30 m telescope during the observational campaign of February, 2015. A first overview of the *NIKA* Crab polarization observations, focusing on instrumental characterization of the polarization system, was given in [Ritacco et al. \(2016\)](#). In this paper we go a step further in the analysis by combining *NIKA* observations with published values at other frequencies to trace the polarized SED of the Crab nebula. We use polarization observations from the *WMAP* satellite at 23, 33, 41, 61 and 94 GHz ([Weiland et al. 2011](#)), from the *Planck* satellite at 30, 44, 70, 100, 143, 217, 353 GHz and from XPOL/30m at 90 GHz ([Aumont et al. 2010](#)).

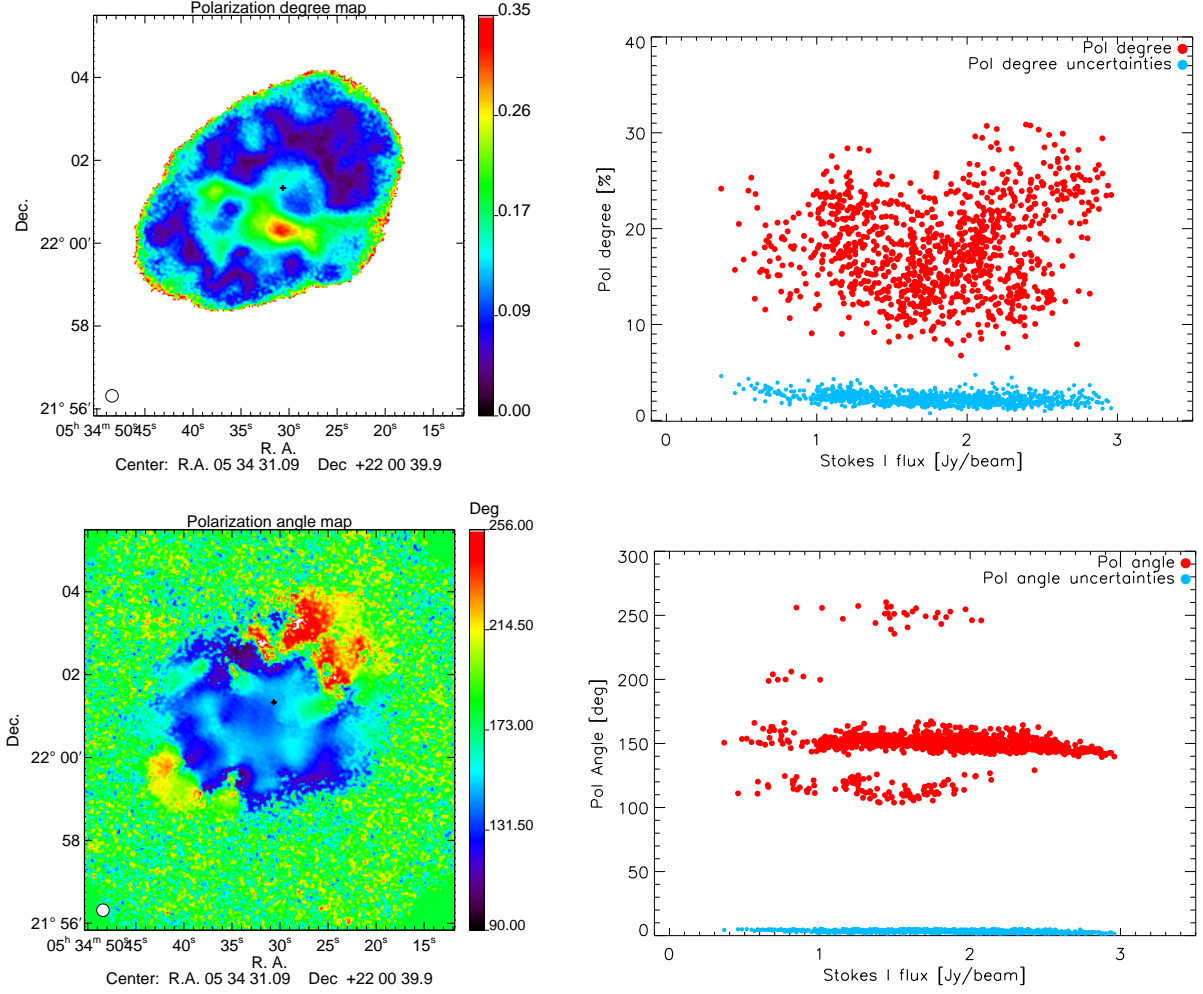
The paper is organized as follows: in Sec. 2 the intensity and polarization maps obtained with the *NIKA* camera are presented together with the polarization degree and angle spatial distributions; Sec. 3 presents the reconstruction of the polarization properties of the Crab nebula in well defined regions; Sec. 4 presents the Crab nebula SED in temperature and polarization; in Sec. 5 we present our conclusions.

## 2. NIKA observations of the Crab Nebula

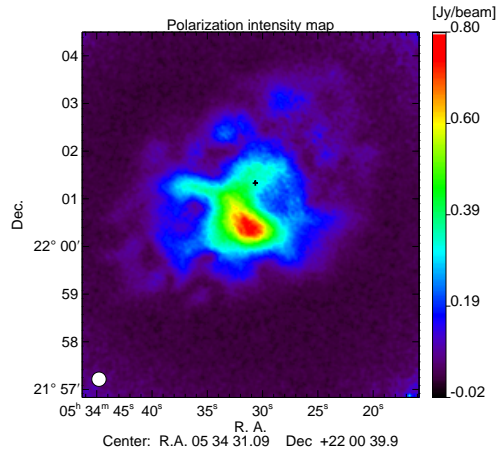
### 2.1. NIKA camera and polarization setup

*NIKA* is a dual band camera observing the sky in intensity and polarization at 150 and 260 GHz with 18 arcsec and 12 arcsec FWHM resolution, respectively. It has a Field-of-View (FoV) of  $1.8'$  at both wavelengths. It was operated at the IRAM 30 m telescope between 2012 and 2015. A detailed description of the *NIKA* camera can be found in [Monfardini et al. \(2010, 2011\)](#) and [Catalano et al. \(2014\)](#).

In addition to having performed intensity observations, *NIKA* was also a test bench for the polarization system of the final instrument *NIKA2* ([Calvo et al. 2016](#); [Catalano et al. 2016](#)), which was installed at the telescope in October, 2015. The po-



**Figure 2.** *Top:* The left panel shows the polarization degree map  $p$ , uncorrected for noise bias, of the Crab nebula. The right panel shows the noise bias corrected  $p$  values as a function of total intensity map (Stokes  $I$ ). The condition  $I_{pol} > 5\sigma_{I_{pol}}$  is satisfied for those values. *Bottom:* on the left we present the polarization angle map  $\psi$  (Equatorial coordinates system) of the Crab nebula. On the right panel the distribution of  $\psi$  values is represented as a function of the total intensity in the case of very high S/N ratio where  $I_{pol} > 5\sigma_{I_{pol}}$ . The cyan dots represent the uncertainties calculated as the dispersion between different observational scans. The black cross marks the pulsar position on the maps.



**Figure 3.** NIKA polarized intensity map of the Crab nebula at 150 GHz. The map shows high polarized emission reaching a value of  $0.8 \text{ Jy beam}^{-1}$ . The telescope beam FWHM is shown in the lower left. The black cross marks the pulsar position.

larization setup of *NIKA* and *NIKA2* consists in a continuously rotating a metal mesh half wave plate (HWP) followed by an analyzer, both at room temperature (for *NIKA*) and placed in front of the entrance window of the cryostat. The Lumped Elements Kinetic Inductance Detectors (LEKIDs) are not intrinsically sensitive to polarization. The HWP rotates at  $2.98 \text{ Hz}$  which modulates the polarization signal at  $4 \times 2.98 \text{ Hz}$ . With a typical telescope scanning speed of  $26.23 \text{ arcsec/s}$ , this provides a quasi-simultaneous measure of Stokes parameters  $I$ ,  $Q$  and  $U$  per beam and places the polarized power in the frequency domain far from the low frequency electronic noise and the atmospheric fluctuations. Ritacco, A. et al. (2017) gives more details on the *NIKA* polarization capabilities and describes the performance of the instrument at the telescope. In particular the sensitivity of the *NIKA* camera in polarization mode was estimated to be  $50 \text{ mJy.s}^{1/2}$  at  $150 \text{ GHz}$ .

*NIKA* has provided the first polarization observations performed with Kinetic Inductance Detectors, confirming that KIDs are a suitable detector technology for the development of the next generation of CMB experiments.



## 2.2. NIKA observations

Crab nebula polarization observations with the *NIKA* camera were performed at the IRAM 30 m telescope in February, 2015. The average effective opacity at 150 GHz was  $\tau = 0.2$ . Fig. 1 shows the Stokes  $I$ ,  $Q$  and  $U$  maps obtained by a co-addition of 16 maps of  $8 \times 6$  arcminutes for a total observation time of  $\sim 2.7$  hours. The maps were performed in equatorial coordinates in four different scan directions:  $0^\circ$ ,  $90^\circ$ ,  $120^\circ$ ,  $150^\circ$ . This allowed us to have the best coverage of the source.

To obtain the Stokes  $I$ ,  $Q$ , and  $U$  Crab nebula maps in Equatorial coordinates, we have used a dedicated polarization data reduction pipeline (Ritacco, A. et al. 2017), which is an extension of the intensity *NIKA* pipeline (Catalano et al. 2014; Adam et al. 2014). The main steps of the polarization pipeline are summarized below:

1. Subtraction of the HWP induced parasitic signal, which is modulated at harmonics of the HWP rotation frequency and represents the most tedious noise contributing to the polarized signal.
2. Reconstruction of the Stokes  $I$ ,  $Q$  and  $U$  time ordered information (TOI) from the raw modulated data. This is achieved using a demodulation procedure consisting in a lock-in around the fourth harmonic of the HWP rotation frequency, where the polarization signal is located.
3. Subtraction of the atmospheric emission in the demodulated TOIs using decorrelation algorithms. In polarization, the HWP modulation reduces significantly the atmospheric contamination **and there is no need to further decorrelate the  $Q$  and  $U$  TOI's from residual atmosphere**. By contrast, in intensity the atmospheric emission fully dominates the signal and to recover the large angular scales we use the 260 GHz band as an atmosphere dominated band like in Adam et al. (2014). **This decorrelation impacts the reconstructed Stokes maps via a transfer function. We have estimated this function with simulated observations of diffuse emission that were passed through the data reduction pipeline, with the exact same scanning, sample flagging and data processing as real data. We found that the power spectrum of the output maps match that of the input one to better than 1% (resp. 5%) on scales smaller (resp. larger) than  $\sim 1'$ . The impact of the data processing is thus negligible compared to uncertainties on absolute calibration on small scales, and its moderate effect on large angular scales is further reduced with the subtraction of a zero level for the photometry (see below). In the following, we therefore neglect this transfer function.**
4. Correction of the intensity-to-polarization-leakage-effect, which was identified in observations of unpolarized sources like the planet Uranus. For point sources the effect was about 3% peak-to-peak, while for extended sources like the Crab nebula it has been found to be the order of 0.5 % peak-to-peak. Ritacco, A. et al. (2017) describes the algorithm of leakage correction developed specifically for *NIKA* polarization observations. Applying this algorithm to Uranus observations the instrumental polarization is reduced to 0.6% of the total intensity  $I$ .
5. Projection of the demodulated and decorrelated Stokes  $I$ ,  $Q$ , and  $U$  TOIs into Stokes  $I$ ,  $Q$  and  $U$  equatorial coordinates maps.

## 2.3. Crab polarization properties

In this section we discuss the polarization properties of the source in terms of polarization degree  $p$  and angle  $\psi$ , which are defined through the Stokes parameters  $I$ ,  $Q$ , and  $U$  as follows:

$$p = \frac{\sqrt{Q^2 + U^2}}{I}$$

and

$$\psi = \frac{1}{2} \arctan \frac{U}{Q}. \quad (1)$$

These definitions are not linear in  $I$ ,  $Q$  and  $U$  and therefore, the observational uncertainties have to be carefully considered, *i.e.*  $p$ ,  $\psi$  are noise biased. Simmons et al. (1980); Simmons & Stewart (1985); Montier et al. (2015) proposed analytical solutions to correct for this bias. For intermediate and high S/N ratio the polarization degree and its uncertainty read:

$$p = \frac{\sqrt{Q^2 + U^2 - \sigma_Q^2 - \sigma_U^2}}{I},$$

$$\sigma_p = \frac{\sqrt{Q^2 \sigma_Q^2 + U^2 \sigma_U^2 + p^4 I^2 \sigma_I^2}}{p I^2}. \quad (2)$$

Furthermore, the polarization angle in a high S/N regime can be approximated by Eq. 1 with the uncertainty

$$\sigma_\psi = \frac{\sqrt{Q^2 \sigma_Q^2 + U^2 \sigma_U^2}}{2(pI)^2}. \quad (3)$$

The angular distribution map of the polarization degree  $p$  of the Crab nebula without noise bias correction is presented on the top left panel of Fig. 2. Tab. 2 presents the total flux  $I$ , polarization degree  $p$  and angle  $\psi$  computed on the pulsar position, highlighted by a black cross on maps, and the Stokes  $I$  peak position. The polarization degree  $p$  reaches a value of  $22.5 \pm 0.7\%$  at the peak of the total intensity, which is consistent with what we observe on the top right panel of Fig. 2, where the variation of  $p$  as a function of the Stokes  $I$  is shown. Here the  $p$  values have been noise bias corrected and satisfy the condition  $I_{pol} = \sqrt{Q^2 + U^2} > 5 \sigma_{I_{pol}}$ . The distribution of the polarization degree appears highly dispersed around a mean value of 20%. The compatibility between the  $p$  value computed at the Stokes  $I$  peak position and the above mentioned plot is expected because we use a very high S/N ratio threshold of  $5 \sigma_{I_{pol}}$ , which restricts the measurement to a small region around the peak of the source. In addition, the  $p$  value found is also consistent within the error bars with POLKA experiment measurement (Wiesemeyer et al. 2014) (cf. Table 2).

The polarization degree decreases towards the edges of the source. Furthermore, the high polarization degree observed at the extremities of the source is misleading and caused by the very low S/N ratio of the Stokes  $I$  map observed in these regions. The variation of  $p$  highlights the interest of high resolution polarization observations of the Crab nebula.

The bottom left panel of Fig. 2 shows the spatial distribution of polarization angle  $\psi$ . We observe a relatively constant polarization angle of about  $150^\circ$  represented here in equatorial coordinates, except for the northern region where the averaged angle is around  $220^\circ$ , and some inner regions with lower polarization angle. These values are confirmed by the bottom right panel that shows the polarization angle distribution as a function of total intensity satisfying the condition  $I_{pol} > 5 \sigma_{I_{pol}}$ .

The sudden change of polarization angle on the northern region was already observed by the XPOL experiment at 90 GHz (Aumont et al. 2010). This together with the variation of the polarization fraction discussed above confirms the need of high angular resolution observations at low and high frequencies for a good understanding of the Crab polarized emission properties. High resolution observations gives the possibility to compare to all CMB experiments as there are significant polarization fraction variations across the source.

We present in Fig. 3 the 150 GHz Crab polarization intensity map  $I_{pol}$  uncorrected for noise bias. We observe a peak at 0.8 Jy beam<sup>-1</sup> and the polarization decreases towards the edges of the nebula.

### 3. Total intensity and polarization fluxes

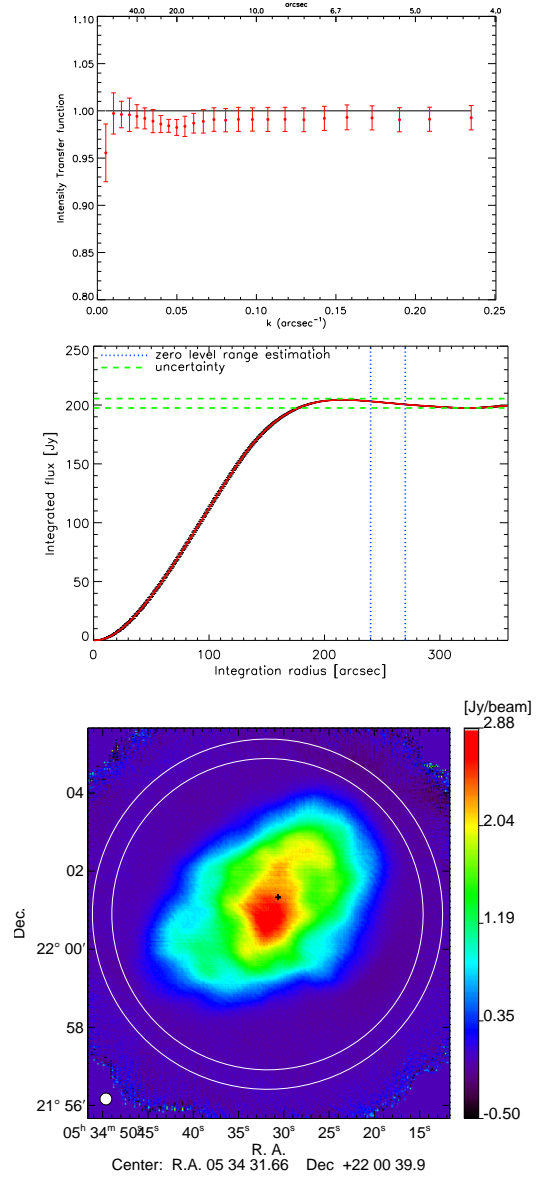
We compute the total flux across the Crab nebula, which has an extent of about 5'x7' as shown in Fig. 1. We use standard aperture photometry techniques to calculate the flux as shown in Fig. 4. We use as center position the center of the map. A zero level in the map, calculated as the mean of the signal measured on an external annular ring region (see bottom panel of Fig. 4) of radius  $4' < R < 4.5'$ , has been subtracted from the map. The total signal estimated is  $204.4 \pm 7.9 \pm 10.2$  Jy. The first uncertainty term accounts for statistical uncertainties computed from fluctuations of the signal at large radii.

We use Uranus for absolute point source flux calibration. The flux of the planet is estimated from a frequency dependent model of the planet brightness temperature as described in Moreno (2010). This model is integrated over the NIKA band-passes for each channel, and it is assumed to be accurate at the 5% level. The final absolute calibration factor is obtained by fitting the amplitude of a Gaussian function of fixed angular size on the reconstructed maps of Uranus, which represents the main beam. For the polarization observational campaign of February 2015 this uncertainty is estimated to be 5% for the NIKA 2.05 mm channel (150 GHz) (Ritacco, A. et al. 2017). Nevertheless, as described in Adam et al. (2014); Catalano et al. (2014), by integrating the Uranus flux up to 100 arcsec, we observe that the total solid angle covered by the beam is larger than the Gaussian best-fit of the main beam by a factor of 28%. As consequence we account for this factor in the estimation of the fluxes. Moreover, Adam et al. (2014) estimated the uncertainty on the solid angle of the main beam to be 4%. Finally, the overall calibration error is estimated to be about 10%, by considering also the uncertainties on the side lobes.

The polarization efficiency factor estimated across the NIKA 2.05 mm spectral band and reported in Ritacco, A. et al. (2017) is  $\rho_{pol} = 0.9941 \pm 0.0002$ . This very small efficiency loss of 0.6 % has a very small impact on the estimation of the polarization fluxes and the calibration error itself.

In order to compare our results with low angular resolution CMB experiments, we present in Tab. 1 the polarization degree  $p$  and angle  $\psi$  integrated values obtained in well defined regions: 5', 7' FWHM from the center of the maps, at the pulsar position and Stokes  $I$  peak position.

The polarization angles are here presented in Galactic coordinates to ease the comparison with the *Planck* (Planck Collaboration et al. 2016a) and *WMAP* CMB experiments (Weiland et al. 2011) and also in Equatorial coordinates within parenthesis. As discussed in Ritacco, A. et al. (2017) a 1.8° uncertainty must be considered because the mechanical motor, in which the HWP is mounted, completes 100 steps per tour of the



**Figure 4.** Top: Transfer function of the intensity data processing. Middle: Cumulative flux of the Crab nebula (top) obtained at 150 GHz over 4' from the center obtained by aperture photometry. The flux has been corrected by a zero level in the map, which corresponds to the mean of the signal calculated in an annular ring as indicated by the white circles on the map (bottom) and by the blue dotted lines on the top. The green dotted line represents the uncertainties measured at large radii.

HWP. Therefore, it is the precision associated to the determination of the HWP zero, corresponding to its optical axis in the cabin reference frame.

Notice that the *Planck* satellite FWHMs are: 33', 24', 14', 10', 7.1', 5.5', 5' arcminutes at 30, 44, 70, 100, 143, 217, 353 GHz, respectively. *WMAP* satellite has FWHMs: 0.93°, 0.68°, 0.53°, 0.35°, <0.23° at 22 GHz, 30 GHz, 40 GHz, 60 GHz, 90 GHz respectively.

As stated in 2.3, we use simple Gaussian estimators in this work that are unbiased only at high S/N ratio. This adds an additional uncertainty on the determination of the degree of polarization  $p$  and on the angle  $\psi$ . While we can delimit regions of the map as a function of their S/N ratios to give

reliable results per pixel, the determination of average quantities over the entire map is more delicate. This is what we must do though to simulate what an experiment with an arcmin beam would see. For example, taking the entire map and averaging it over Gaussian beams of  $5'$  and  $7'$  FWHM, we find  $-83.6912 \pm 0.0004^\circ$  and  $\psi = -83.6491 \pm 0.0004^\circ$  respectively, while we find  $-88.15 \pm 0.8$  if we consider only pixels with a  $I_{pol}$  S/N larger than XXX. For  $p$ , using the same selection, we find YYY...

Fig. 5 shows the polarization fraction (top) and polarization angle (bottom) of the Crab nebula as a function of the frequency as measured by five different instruments: *Planck* (Planck Collaboration et al. 2016a), *WMAP* (Weiland et al. 2011), *XPOL* (Aumont et al. 2010) and *NIKA* (this paper).

The solid line in both figures represents the weighted-average found using all the observations shown and the dashed lines represent the uncertainties estimated on the weighted-average. Considering only low frequency data ( $<200$  GHz) and excluding the *XPOL* data (see below) we find that the degree of polarization of the Crab nebula at arcmin scales is  $7.1 \pm 0.1\%$ . In terms of degree of polarization most data sets are consistent at the  $2\sigma$  level with the weighted-average value. For *XPOL* the discrepancy can probably be explained by the lower sensitivity of the single channel *XPOL* experiment to the lower than average polarization of the outer parts of the nebula.

In the case of *Planck* the observed discrepancy at high frequency could be due to an evolution of the Crab nebula polarization properties with frequency, but this hypothesis is disfavored by the polarization angle measurements that seem to be consistent with the mean value across all frequencies.

The weighted-average of the polarization angle of the Crab nebula at  $\geq 5'$  scales is therefore estimated considering all available observations, is  $-88.1^\circ \pm 0.3$ . All the observations shown on the bottom panel of Fig. 1 agree within  $1\sigma$  with this value except for the *Planck* value at 30 GHz, which is slightly high.

## 4. Crab Spectral Energy Distribution in intensity and polarization

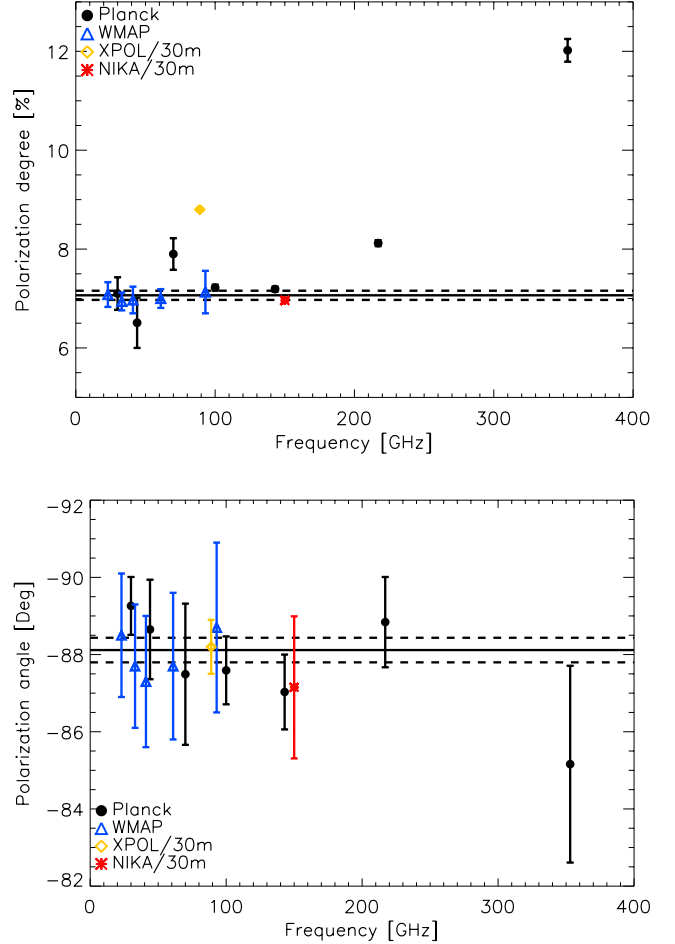
### 4.1. Intensity

The total intensity emission of the Crab nebula at radio and millimeter wavelengths (from 1 to 500 GHz) is mainly due to synchrotron emission and can be well described by a single power law of the form:

$$I_\nu = A(\nu/1\text{GHz})^\beta \quad (4)$$

with spectral index  $\beta = -0.296 \pm 0.06$  (Baars et al. 1977; Macías-Pérez et al. 2010). Further, the Crab nebula is fading with time at a rate of  $\alpha = 0.167 \pm 0.015\% \text{ yr}^{-1}$  (Aller & Reynolds 1985). These results suggest a low frequency emission produced by particles accelerated by the same magnetic field. Macías-Pérez et al. (2010) have shown also that there is no evidence for an extra synchrotron component and for a component of thermal dust emission at these wavelengths. The direction of the polarization is therefore expected to be constant across frequency range 30-300 GHz while the polarization degree may vary.

We show in Fig. 6 the flux of the Crab nebula as a function of frequency. The fluxes in the radio domain were taken from Dmitrenko et al. (1970) and Vinogradova et al. (1971). We also show microwave and mm wavelengths fluxes from *Archeops* (Macías-Pérez et al. 2007), *Planck* (Planck Collaboration et al. 2016a), *WMAP* (Weiland et al. 2011), *XPOL/30m* (Aumont et al. 2010), *MAMBO/30m* (Bandiera



**Figure 5.** *Top:* polarization degree as a function of frequency as measured by *Planck* (black dots), *WMAP* (blue triangles), *XPOL* (green diamond) and *NIKA* (red crosses). The *NIKA* value has been estimated by integrating in a radius of  $5'$  as given by *XPOL* (Aumont et al. 2010). The solid line represents the weighted-averaged degree of polarization accounting for low frequency values ( $<200$  GHz) and excluding *XPOL*. Dashed lines are the uncertainties. *Bottom:* polarization angles in Galactic coordinates for the same five experiments given above. The solid line represents the weighted-averaged polarization angles. Notice that the *NIKA* value satisfies the condition  $I_{pol} > 3\sigma_{I_{pol}}$  to avoid noise bias.

et al. 2002) and *GISMO/30m* (Arendt et al. 2011). The measured *NIKA* total intensity flux at 150 GHz is shown in red. We present in the figure the best-fit model obtained assuming a power law as given by Eq. 4.1 to the data computed by  $\chi^2$ -minimization using observations below 100 GHz (cyan line). The amplitude and spectral index found are:

$$A = 980.6 \pm 0.7, \quad \beta = -0.3151 \pm 0.0002. \quad (5)$$

For illustration we also show the best-fit model to the data found by Macías-Pérez et al. (2010) in a previous analysis (green line). Notice that fading of the source is accounted for in the best-fit models as well as in the data presented in the figure. The estimated best-fit model from this paper is slightly different with respect to the previous analysis mainly due to the addition of re-



	$I$ [Jy]	$I_{pol}$ [Jy]	$p$ [%]	$\psi$ [°]	$I_{pol} > 3\sigma_{I_{pol}}$
7' FWHM centered	207.9±0.3	14.4±0.1	6.93±0.06	-83.4 (141.02)±0.2	-87.14 (144.8)±0.07±1.8*
5' FWHM centered	207.2±0.1	14.4 ±0.04	6.96±0.02	-82.3 (139.9)±0.7	-87.15 (144.8)±0.04±1.8*

**Table 1.** \*A systematic angle uncertainty of 1.8° must be considered in the polarization angle error budget. Total flux  $I$ , polarized intensity flux  $I_{pol}$ , polarization degree  $p$ , and angle  $\psi$ . The values have been calculated within 7' and 5' radius from the center of the map. The polarization angle has been also calculated where  $I_{pol} > 3\sigma_{I_{pol}}$  to avoid the contribution of pixels biased by the noise. The polarization angle is here given in Galactic coordinates and in Equatorial coordinates within parenthesis. An total calibration error of 10 % must be accounted for and propagated to the polarization estimates.

	Pulsar		
	$I$ [Jy]	$p$ [%]	$\psi$ [°]
NIKA	0.97 ± 0.01	13.8 ± 0.7	151.4 ± 1.5 ± 1.8 (-93.8)
POLKA	1.63	25.3 ± 3.0	145.1 ± 3.3
XPOL		13.9 ± 0.6	158.1 ± 0.5
SCUPOL		14.3 ± 1.8	140.0 ± 2.8
	Peak		
NIKA	1.14 ± 0.01	22.5 ± 0.7	145.8 ± 0.8 ± 1.8 (-88.15)
POLKA	1.72	25.0 ± 3.1	151.7 ± 3.5
XPOL		25	149.0 ± 1.4
SCUPOL		18.7 ± 1.5	146.1 ± 2.1

**Table 2.** Total flux  $I$ , polarization degree  $p$ , and angle  $\psi$  estimated at the pulsar position evidenced on the maps by a black cross and at the peak of the Stokes  $I$  map. The polarization angle is here given in Equatorial coordinates and in Galactic coordinates within parenthesis. Values estimated by POLKA, XPOL, SCUPOL reported in Wiesemeyer et al. (2014) are also represented for comparison.

cently published results by *Planck* (Planck Collaboration et al. 2016a) and *WMAP* (Weiland et al. 2011). However, the result found at 150 GHz by the *NIKA* camera agrees at 1  $\sigma$  level with both models. Notice that, as discussed before, the XPOL/30m integrated flux is low with respect to expectations from both power law models.

The *Planck* data at 100, 143, 217 and 353 GHz are consistent with a much steeper SED with respect to the lower frequency data suggesting an extra emission component due to a different population of relativistic electrons (Ginzburg & Syrovatskii 1965; Gomez et al. 2012). *Archeops*, MAMBO/30m, GISMO/30m and *NIKA*/30m do not seem to support such a low frequency break in the Crab nebula SED. Moreover, Macías-Pérez et al. (2010) found a break for frequencies above 1000 GHz. However, considering the break shown by *Planck* satellite data we assume a simple two-power-law model<sup>1</sup>:

$$I_\nu = A(\nu/1\text{GHz})^\beta \quad \text{for } \nu < 100\text{GHz} \quad (6)$$

and

$$I_\nu = A_H(\nu/1\text{GHz})^{\beta_H} \quad \text{for } \nu > 100\text{GHz} \quad (7)$$

By  $\chi^2$ -minimization we find  $A_H = 8.6 \pm 0.45$  and  $\beta_H = -0.71 \pm 0.09$ . We note that *NIKA* data is consistent with this model at the 2 $\sigma$  level. More data at mm wavelengths are needed to better understand the observed break of the SED at about 90 GHz.

#### 4.2. Polarization

The total intensity of the Crab nebula has been monitored over decades across a large range of frequencies. By contrast the amount of polarization observations is really small. Recent results provided by *Planck* (Planck Collaboration et al. 2016a),

*WMAP* (Weiland et al. 2011) and XPOL (Aumont et al. 2010) together with *NIKA* allow us to trace the spectrum of the polarized flux as shown in Fig. 7. The *Planck* data points show an oscillating behavior while the *WMAP* ones are more consistent with a power law model. This figure show a tighter correlation of the XPOL measurement with other data because it is not affected by the Stokes  $I$  quality. Furthermore, the average slope of the SED when approached by a single power-law seems to be different for the *Planck* and *WMAP* data. To investigate this issue further we assume a power law model for the data as in Eq. 4.1, and fit separately the *Planck* and *WMAP* data. We restrict the fit to the data with  $\nu < 100$  GHz. We show in Fig. 7 in black and blue the best-fit models for the *Planck* and *WMAP* data, respectively.

Accounting for fading the fitted amplitudes and spectral indexes are:

- best-fit parameters for *WMAP* only data:

$$A = 78.9 \pm 7.8 \quad , \quad \beta_p = -0.35 \pm 0.03; \quad (8)$$

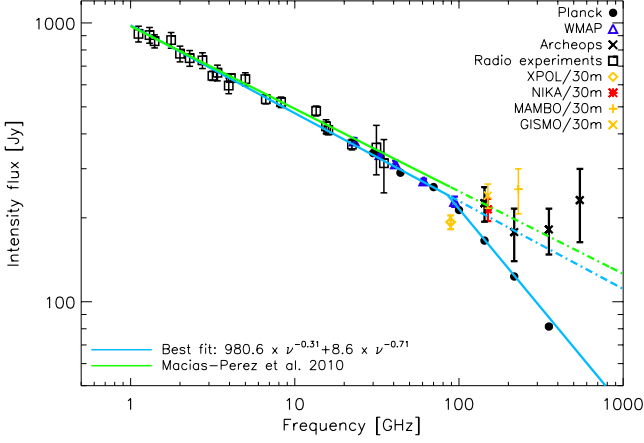
- best-fit parameters for *Planck* only data:

$$A = 179.1 \pm 15.4 \quad , \quad \beta_p = -0.54 \pm 0.02. \quad (9)$$

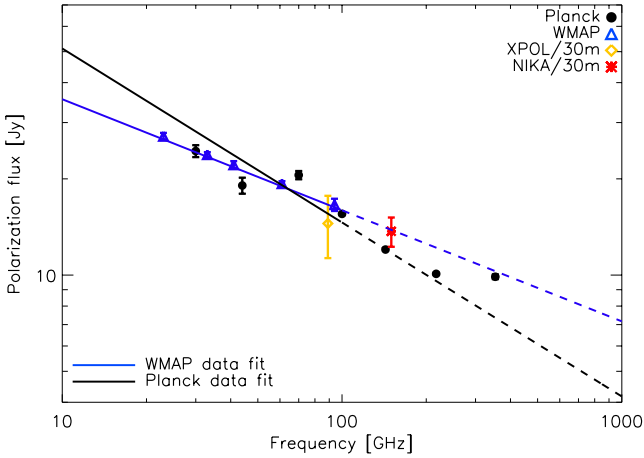
The *WMAP* best-fit model is in agreement with *Planck* observations at 30, 100 and 353 GHz. The *Planck* best-fit model seems to overestimate the polarization flux at low frequency. The *NIKA* data at 150 GHz are consistent at the 1 $\sigma$  level with both the *WMAP* and *Planck* best-fit models. In contrast to the intensity data, here a 100 GHz break is less conspicuous. We find no clear hint of a break in the polarization SED at frequencies above 90 GHz, hence putting pressure on the hypothesis of multiple electron populations to explain the SED break in intensity.

We have also estimated the spectral index of the Crab nebula polarization emission at high frequency using the map obtained by SCUPOL (Matthews et al. 2009) at 352 GHz (850

<sup>1</sup> A detailed physical description of the SED of the Crab nebula is out of the scope of this paper.



**Figure 6.** Crab nebula total intensity SED as obtained from *Planck* (Planck Collaboration et al. 2016a), *WMAP* (Weiland et al. 2011), *Archeops* (Macías-Pérez et al. 2007), radio experiments (Dmitrenko et al. 1970; Vinogradova et al. 1971), *XPOL/30m* (Aumont et al. 2010), *NIKA30m* (this paper), *MAMBO/30m* (Bandiera et al. 2002) and *GISMO/30m* (Arendt et al. 2011) data. The green line shows the model derived by a previous analysis discussed in (Macías-Pérez et al. 2010). The best-fit obtained by the analysis in this paper is shown in cyan line. Both, the best-fit models and the data account for the Crab nebula fading with the time.



**Figure 7.** Crab nebula polarization flux SED as obtained from the *Planck* (Planck Collaboration et al. 2016a), *WMAP* (Weiland et al. 2011) and *NIKA* (this paper) data. The two best-fit models presented have been estimated using only *WMAP* data (blue line) or *Planck* data only (black line).

$\mu\text{m}$ ) and the *NIKA* map. Considering only the region observed by SCUPOL we obtain:

$$\beta_p = -0.33 \pm 0.01. \quad (10)$$

This result is in good agreement with the *WMAP* best-fit model spectral index.

## 5. Conclusions

The Crab nebula is considered as a celestial standard calibrator for CMB experiments in terms of polarization degree and angle. An absolute calibration is particularly important for the measurement of the CMB polarization B-modes, which are a window towards the physics of the early Universe.

We have reported in this paper first high angular resolution polarization observations of the Crab nebula at 150 GHz with the *NIKA* camera. These observations have allowed us to accurately map the spatial distribution of the polarization fraction and angle. We find an averaged polarization angle of  $-87.15 \pm 0.04 \pm 1.8$  within a region of  $5'$  radius with respect to the central coordinates of the map R.A.:  $05^h 34^m 31.09$  s, Dec.:  $+22^\circ 00' 39.9''$  (J2000). This is consistent with previous measurements from *Planck*, *WMAP*, and *XPOL/30m*. Using all available polarization data to date we conclude that the polarization angle of the Crab nebula is consistent with being constant with frequency, from 23 GHz to 217 GHz, at arcmin scales with a value of  $-88.1^\circ \pm 0.3$  in Galactic coordinates. In addition, there is a strong case for a constant polarization degree below 200 GHz. Above this frequency, *Planck* results are not consistent to each other.

Moreover, we have characterized the intensity and polarization SED of the Crab nebula. In intensity, recent *Planck* data show a break in the SED that could be explained by the existence of different populations of relativistic electrons. However, they are inconsistent with other data. In polarization, we find that the data are overall consistent with a single power law spectrum as expected from synchrotron emission. However, we find some discrepancies between data sets, which will require further mm measurements at high resolution for a better understanding of the physics at play. Among future polarization experiments, *NIKA2* (Calvo et al. 2016), will provide high sensitive polarization observations of the Crab nebula adding a 260 GHz map at  $11''$  resolution.

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