

ASTR 400B Research Project: The Fate of Sun-like Stars in M31's Disk

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1. INTRODUCTION

Gravity is a far reaching force, allowing objects that are even hundreds of millions of light-years away from each other to interact and influence each other's trajectories through space. In the case of the galaxy we inhabit, the Milky Way, it is gravitationally interacting with other nearby (within several million ly) galaxies, the largest of which is the Andromeda Galaxy (M31). Both our Milky Way and Andromeda are spiral galaxies consisting of two main stellar components, the disk and the bulge, and notable spiral arms made of gas and dust. The disk consists of typically younger stars that orbit the galactic center in a wide, circular, and thin distribution, and it is the component of particular interest in this report. The bulge is an older, puffier distribution of stars located close to the galactic center. The Milky Way and M31 also both have several, smaller galaxies orbiting them, the most massive of which are the Large Magellanic Cloud and the Triangulum Galaxy (M33) respectively. These less massive galaxies are gravitationally bound to either the Milky Way or M31, meaning they lie within their host galaxy's gravitational potential well and do not have enough kinetic energy to escape this well, and are known as satellite galaxies. The Milky Way, Andromeda, and their satellite galaxies form what is known as the Local Group. Close interactions between galaxies (such as the Milky Way and M31) can involve moving billions of solar masses, which in turn takes equally enormous amounts of energy. Potential energy is converted into kinetic energy as two objects come together, and in the case of two galaxies, some of this kinetic energy is imparted onto the stars of either system. The more massive a system is, the deeper a gravitational potential is, and the more energy is required to remove something from that potential, so while the stars in these galaxies may have higher velocities than before the interaction occurred, they still may not have enough to entirely escape the system due to the increased potential of the combined masses of the two galaxies.

A galaxy is a collection of gravitationally bound stars orbiting a common center of mass and whose rotation curve cannot be explained by the collective observed baryonic mass (Willman & Strader 2012). In other words, there is a 'missing' mass distribution that we cannot observe in any wavelength of electromagnetic radiation that would cause the flattened rotation curves we observe. For the purposes of this report, this mass distribution is known as a dark matter halo and the mass of this halo is the dominant component of a galaxy's gravitational potential well. Despite their scale, the structures of galaxies are not constant over time, however. In example, star-forming galaxies eventually use up the dust and gas within them, and gas-rich features such as spiral arms may dwindle, while the galaxy's light will become redder as stars age and metallicity will increase as massive stars die. Such changes in a galaxy's structure and composition are known as galaxy evolution. If one galaxy gravitationally interacts with another, especially in the cases of near-flybys and mergers, we are able to see a dramatic form of galaxy evolution occur. At this point in the study of galactic interactions, astronomers are well aware that the two most massive galaxies in our Local Group, the Milky Way and the Andromeda Galaxy, are on a collision course. Now, we can try to understand how individual components of these galaxies interact and what sorts of remnants we would expect to see using simulations that behave as though we could observe our galaxies over billions of years with ever increasing resolution. Within these components, we can also consider how we expect individual particles to behave before, during, and after the merger, and how the distribution of these particles (especially the materials the contain) may dictate the merger's future.

In the third paper of The M31 Velocity Vector series, probabilities of the outcome of our Sun's position are detailed by considering "candidate suns". These "suns" are Milky Way disk particles that are within 10 percent of the Sun's distance to the center of the Milky Way, have a circular planar velocity within ten percent of the Sun's circular planar velocity, and an out-of-plane velocity less than 30 km/s. This collection was selected at $t = 3$ billion years into the simulation, and allowed for just under 9000 sample suns, whose positions and velocities could be determined at the

simulation’s end, 10 billion years in the future. In all cases, the solar candidates remained bound, though at varying distances from the center of the MW/M31 merger, with 85 percent of candidates located at radii greater than 8.29 kpc (van der Marel et al. 2012). The positions of stars as a result of the MW/M31 merger have been thoroughly

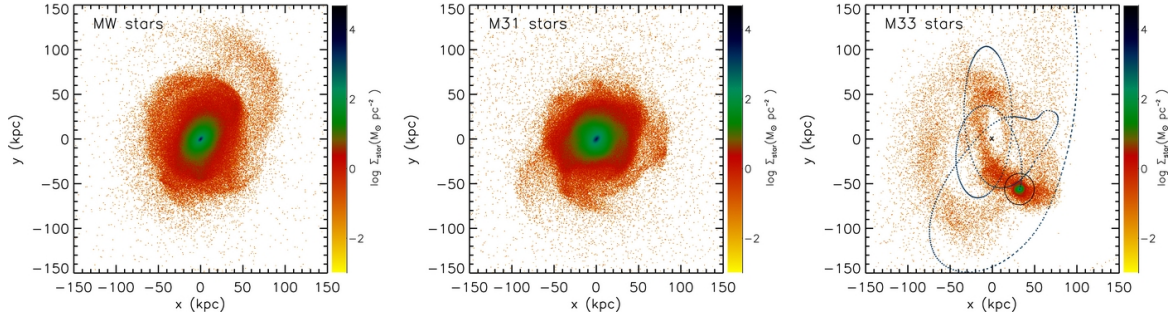


Figure 1. Fig. 6 from van der Marel et al. (2012) visualizing the positions of stellar particles in (L to R) the MW, M31, and M33 at $t = 10$ Gyr, with the color scale representing surface mass density. The MW and M31 maintain loosely circular central distributions of about $10 M_{\odot}/\text{kpc}$ out to about 50 kpc from the COM of each (shared by this point in the simulation), with thinner stellar streams extending out at least another 100 kpc. The stars in M33 have been widely distributed into streams, though a ‘central’ concentration can still be distinguished. This figure clearly demonstrates how stars can be distributed as the MW/M31 merger relaxes after being formed, but it does not tell us where these stars originated in their parent galaxy.

When considering stars involved in a merger of two galaxies, there are many facets of their spacial and kinematic properties that can be analyzed. Adding a third, though smaller, galaxy into this situation allows for further study. In particular, we can compare the kinematics of stars in a large galaxy with a satellite (M31 in this simulation) to those in a galaxy without such a companion (the MW). Additionally, we can compare the initial and final distribution of stars in any of the three galaxies. Not only this, but we could describe where stars of any specifically chosen mass, position, or velocity characteristics, and learn about how their trajectories differ between the two massive galaxies. Investigations about the Sun’s fate as a result of the MW/M31 merger have been conducted in earlier works, such as in Cox & Loeb (2008), though this work isn’t nearly as high resolution as van der Marel et al. (2012), as only 700 or so candidate suns were considered. As of now there doesn’t appear to be literature specifically dealing with stars at the Sun’s position in either M31 or M33, and as such, the fates of these stars as a result of the MW/M31 merger are generally unexplored.

2. THIS PROJECT

This is not complete, I’m sorry. I’m not expecting passing marks or an extension, but I’ll continue to work on this until it is complete.

This project will aim to create an understanding of the fates of stars in M31 that have initially similar spacial properties as the Sun. To begin, candidate suns will be chosen in a similar manner as in van der Marel et al. (2012), detailed in paragraph 3 of the above Introduction, but for M31 instead of the Milky Way. Specifically, this project will consider stars initially located within 10% of 8.29 kpc away from the center of mass of M31 (final range to be determined) that have stable, near circular orbits at this radius as determined by M31’s rotation curve. These parameters may be changed to ensure that a comparable number of particles is selected (about 9000). Data used to simulate the merger will also be from this paper. This project will focus on how the positions of the candidate suns change over time, and whether or not any of these stars become unbound, as the latter was not observed in van der Marel et al. (2012).

With this project, we will be able to better understand how the dynamics of M31 may differ from those of the Milky Way by comparing the fates of candidate suns between the two galaxies. Another consideration that can be explored is how many candidate suns exist in M31 as have been identified using the same methods of selection as van der Marel et al. (2012) assuming sun-like stars are present around 8.29 kpc from the center of mass of M31 in the same way as in the MW.

not formatted but here are my current thoughts until I can write this better: I want to know how close candidate suns in M31 get to their escape velocity, and where those that get closest are located after the merger (perhaps over the next few billion years). I'll rotate M31 to be edge on so I can easily determine V_{circ} for the sun candidates. Right now I'm planning a histogram of stars vs V/V_{esc} and potentially a 2d plot showing where stars with a velocity close to their V_{esc} (fraction tbd) are with respect to the merger COM.

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2.1. Approach

In order to determine which disk particles will be considered as candidate suns, I will need to read in the data from [van der Marel et al. \(2012\)](#) and create a separate array containing the indices of particles that fit the selection criteria. To do this, I will use code created in Homework 2 (ReadFile) and collect the indices particles within 10 percent of 8.29 kpc. From this smaller list of particles, I will calculate $V_{\text{circ}} (= \sqrt{V_x^2 + V_y^2})$ of each and collect the indices particles within 10 percent of 8.29 kpc that have V_{circ} within 10 percent of 293 km/s and $V_{\text{oop}} (= V_z)$ less than 30 km/s. I will need to create a function that selects particles that fit these criteria, but this can be pieced together using np.where. I will need to save the indices of these particles so I can compare their positions through different time snapshots. I will also need to calculate and compare V_{esc} and $V_{\text{tot}} (= \sqrt{V_x^2 + V_y^2 + V_z^2})$ at the final radial positions for all of the candidate suns to determine if any become unbound from the MW/M31 merger remnant.

We could consider stars much closer in or much farther out from the initial center of mass of M31 than about 8 kpc, but this radius in particular is of interest because it reflects our own Solar System's position in the Milky Way. We also could consider a larger portion of candidate suns, but by narrowing our range of objects of interest down, we can pay better attention to how they behave as the merger is underway, perhaps finding patterns that would have otherwise been obscured had we considered more particles. In trying to understand the details of galaxy mergers, we need to consider the fates of objects at specific positions so we can know how ordered or disordered mergers truly are. In other words, we'll be able to determine whether or not stars around 8 kpc in M31's disk end up in a similar configuration in the combined MW/M31 remnant, and if they will be scattered widely or follow some stream pattern. With this information, we can better understand features we may observe in other mergers that are further along than the Milky Way and M31, as well as understand how the metals found in solar analogs will end up distributed as a result of the MW/M31 merger.

2.2. Hypothesis

We have seen in class that the Andromeda Galaxy and the Milky Way have comparable masses, and thus will have comparable effects on each other. However, M33 is also interacting with M31. We see the effects of the MW/M31 merger on candidate suns in the Milky Way in [van der Marel et al. \(2012\)](#), but these stars are not simulated as being affected by a close satellite galaxy as those in M31 are. Taking this into consideration, I predict that the positions of candidate suns in M31 will display a wider range of positions by 10 Gyr in the future than those in the Milky Way. I also predict that it is unlikely that any of M31's candidate suns will become unbound from the merger remnant because the Milky Way's candidate suns did not become unbound even when tugged by both M31 and M33 at close range.

2.3. Figure

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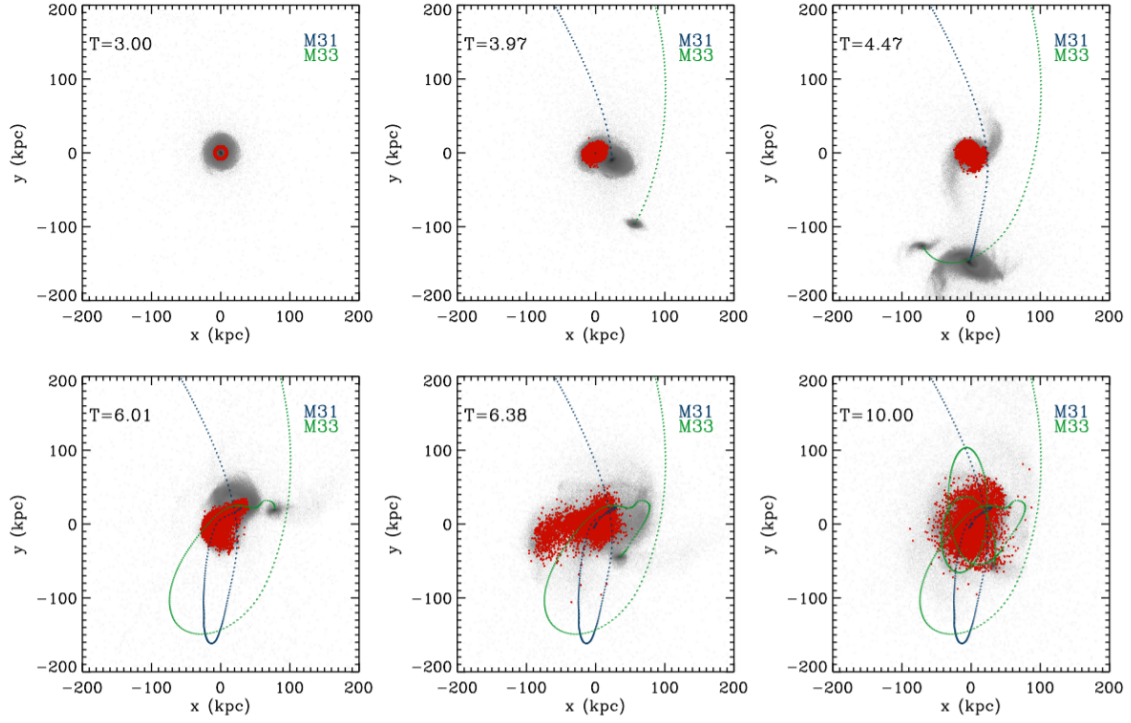


Figure 2. Fig. 5 from [van der Marel et al. \(2012\)](#) visualizing the positions of candidate suns in the disk of the Milky Way (in red) from $t = 3$ Gyr to $t = 10$ Gyr