

Quantum limited parametric amplifier for millimeter and submillimeter astrophysics

Abstract

This proposal is to develop a technology that would significantly improve the noise performance and available instantaneous bandwidth for coherent receivers operating up to ≈ 1 THz, i.e (sub)millimeter wavelengths. We propose to build a Traveling-wave Kinetic Inductance Parametric Amplifier (TKIP) to provide gain before mixer down conversion. The TKIP can provide quantum limited noise performance across an octave bandwidth, which relaxes the requirements on the mixer and afterwards. Although the TKIP could be used in receivers at a variety of facilities, (e.g SOFIA, CCAT), we propose to demonstrate its capabilities in an ALMA band 3 receiver, 80 GHz to 116 GHz. Here, use of the TKIP would mean a factor ≈ 2 better noise and factor ≥ 5 broader instantaneous bandwidth. Such improvement in receiver performance would accelerate pursuit of and provide new access to scientific objectives at these wavelengths.

1 Statement of problem

High resolution spectroscopic observations in millimeter to submillimeter wavelengths are made using telescopes with heterodyne receiver systems. In these receivers, power radiated from the sky is optically focused into a feed horn, then mixed to a lower intermediate frequency (IF) using an SIS mixer, then the IF signal is amplified with a low noise transistor amplifier (LNA), and then further processed and detected. There is a loss in signal power when the mixer performs the conversion to the IF frequencies. Additionally, the LNA adds noise power to the IF. A low power, i.e ‘faint’, astrophysical signal can therefore escape detection because the conversion loss of the mixer shifts the signal power below the noise power of the LNA. This limitation on *sensitivity* to faint sources is one problem to be addressed by this project.

The mixer has a certain bandwidth over which it is sensitive to signal. The lowest noise LNAs provide the required gain over a bandwidth that is much narrower than the mixer bandwidth. Thus, to measure a signal over the entire mixer bandwidth, it is necessary to increment the local oscillator (LO) frequency of the mixer multiple times to shift the signal into the bandwidth of the LNA. Incrementation of the LO comes at the expense of observation time. This limitation on *instantaneous bandwidth* is another problem to be addressed by this project.

2 Proposed solution

We propose to develop an amplifier to be inserted immediately before the mixer in a heterodyne receiver setup. This amplifier would provide enough gain over the full mixer bandwidth to relax the noise performance required of the LNA, with improvement to the receiver's sensitivity. This also permits the use of an LNA that covers the full mixer bandwidth.

Development of this new amplifier, to be called the wTKIP¹, would constitute a major breakthrough in the performance of heterodyne receiver systems for (sub)millimeter waves. As a specific demonstration of its impact, we propose to develop a W-band (80 - 116 GHz) amplifier suitable for use in Atacama Large Millimeter/Submillimeter Array (ALMA) Band 3 because: one, the wTKIP would significantly improve receiver capabilities in this band and therefore has great scientific benefit; two, the group of my proposed JPL advisor has already initiated work on the wTKIP using seed funding, which is now exhausted, provided by the ALMA technology development program; and three, the ALMA team has encouraged the JPL group to apply for additional funding to demonstrate a wTKIP receiver on their array. The proposed work fits nicely between reasons two and three because it aims to develop the TKIP technology to state where it is ready to be used in an ALMA band 3 receiver. After demonstration on ALMA, the goal is to use the technology at other facilities, such as SOFIA, where a TKIP wideband receiver system would enable spectroscopic capabilities

¹Throughout this proposal, "TKIP" refers to those devices for use at unspecified frequency bands, whereas wTKIP refers to the proposed device which is optimized for use at the W-band and μ TKIP refers to the microwave device described in [5] .

at a rate comparable with HIFI on Herschel whose cryogenic mission has now ended; and CCAT, where use of the TKIP would enable a very fast spectral line survey instrument in the atmospheric windows to complement ALMA; and other applications like coherent receiver systems for interferometry or correlation receivers to tackle difficult continuum observations such as the CMB B-mode polarization.

3 Scientific merit

Millimeter and submillimeter wave astronomy can provide information about processes associated with clouds of dust and molecular gas, such as star formation. Progress toward understanding these processes is in accordance with the scientific priorities declared in the National Academies 2010 decadal survey² as well as NASA’s overall scientific goals³. In general, the proposed research will result in vastly improved observing performance and efficiency for accomplishing these priorities.

Molecular gas and dust emission and absorption in submillimeter galaxies Submillimeter galaxies (SMGs) trace a large fraction of the star formation activity in the universe and are dominated by objects undergoing rapid star formation at relatively high redshifts ($1 < z < 3$). ALMA can locate SMGs precisely and obtain a redshift using sequential observations with different local oscillator tunings to cover an entire ALMA band to detect one or more redshifted CO lines. Currently, covering the whole of ALMA band 3 requires setting the local oscillator (LO) to 5 different frequencies. An increase in instantaneous bandwidth to cover the entire band combined with a reduction in system temperature would allow detection of one or more CO lines at all redshifts in a single LO tuning and provide an increase in continuum sensitivity that would enable at least an order of magnitude increase in the speed of imaging and spectroscopy of high redshift SMGs.

²See decadal survey chapter 2 on “the compelling questions for the next decade and beyond” with regard to “origins”.

³See science goals listed in section 4.4 of the 2014 NASA Science Plan.

Galaxy clusters The Sunyaev-Zel’dovich (SZ) effect is a distortion of the cosmic microwave background (CMB) spectrum caused by the scattering of CMB photons as they pass through the hot intracluster medium (ICM) of massive galaxy clusters. The brightness of the spectral distortion is independent of redshift of the cluster. The SZ effect signal is proportional to the integrated pressure of the ICM plasma, which makes it an excellent proxy for cluster mass. These characteristics make the SZ effect a useful window into the astrophysical processes that shape galaxy clusters. Recent observations of clusters of galaxies have shown a large fraction to have an irregular morphology due to merger events and feedback within the clusters. These are evidenced by the presence of shock fronts and cold fronts, as well as structure in these objects. Wide-field high angular resolution imaging with moderate spectral resolution in bands from 15 GHz to 300 GHz would allow unprecedented determination of the cluster gas temperature and density profiles and therefore the cluster mass and dynamics especially when combined with Xray data from e-Rosita and optical weak lensing surveys. These measurements could only be made with a significant increase in sensitivity, frequency coverage and mapping speed over the existing single dish and interferometric instruments.

Atmospheric window	100 GHz
Opacity	0.02
T_{sys} for $T_{\text{rec}} = 5 \frac{h\nu}{k_B}$	30 K
T_{sys} for $T_{\text{rec}} = 1.5 \frac{h\nu}{k_B}$	11.3 K
Integration time reduction	$\times 3.9$

Table 1: Expected performance improvement for wTKIP assuming identical instantaneous receiver bandwidths in the 100 GHz atmospheric window and atmospheric opacity. The opacity assumes 0.66 mm precipitable water vapor, commensurate with a site like Cerro Chajnantor. $5 \frac{h\nu}{k_B}$ represents the current best receivers and $1.5 \frac{h\nu}{k_B}$ is the expectation for the wTKIP.

Star formation in the Milky Way Star formation in the Milky Way takes place in molecular clouds that are of high enough density. Understanding how molecular clouds of low density transition to clouds with density high enough to support star formation and why such a small fraction of the cloud mass is contained in dense gas is key to determining the dominant physical processes that control the

star formation rate.

Spectral lines from different molecular transitions probe different volume densities and temperatures, and time-dependent chemistry favors the production of different molecules at

distinct times. With ≈ 35 GHz of instantaneous bandwidth (the entire W-band), a receiver based on this technology would have the ability to observe multiple molecular transitions simultaneously and probe these different regions. With suitable back end processing power, this spectral multiplex advantage would decrease observing time by the number of lines simultaneously observed, modulo integration time differences due to varying line strengths. Such back end processing bandwidth is achievable at low price using a bank of 12 CASPER ROACH2 spectrometer boards [11], for example.

4 Technical merit

General figures of technical merit of a TKIP include the following:

- Currently state-of-the-art SIS mixers and IF amplifiers limit instantaneous receiver bandwidth ≤ 10 GHz. Use of a TKIP as a front end amplifier would permit the use of lower gain, wider bandwidth mixers and IF amplifiers in the receiver signal path. Not only would it be easier to approach the quantum limit (QL) with such a receiver, but the IF bandwidth could be widened, we believe, by over an order of magnitude⁴ (Fig. 3a), increasing the speed for line surveys by the same factor and greatly improving the sensitivity for continuum observations.
- The TKIP is significantly more sensitive than transistor amplifiers. Theoretically, the TKIP should reach the standard QL ($T_{\text{noise}}k_B = 1/2$ photons per second per Hz). Transistor amplifiers at best add noise about five times higher than this limit.
- Superconducting paramps using Josephson junctions operate at the QL, but these are narrow band devices with fractional bandwidth of at most only a few percent due to the use of resonant circuits. The TKIP has a bandwidth approaching that of transistor amplifiers, with octave bandwidth being readily achievable.
- Other superconducting paramps have operated at tens of GHz and transistor amplifiers

⁴For the case of ALMA, the bandwidth of a wTKIP would completely cover band 3, which accords to a factor ≥ 5 increase in instantaneous bandwidth.

operate with reduced fractional bandwidth up to a few hundreds of GHz. Owing to its simple distributed design, the TKIP will potentially operate up to the gap frequency of its superconducting material; that is, $\lesssim 1.4$ THz or $\gtrsim 214$ μm for NbTiN.

- Compared to transistor amplifiers, the TKIP requires lower temperature operation, around 1 K to 4 K (Fig. 3b). However, one, commercially available pulse tube coolers easily reach that temperature; and two, when the TKIP replaces an SIS mixer, the TKIP can reuse the existing cryogenic infrastructure put in place for the SIS mixer.

For the particular application of the wTKIP on ALMA band 3, Tab. 1 summarizes the expected performance improvements. In addition to the shortened integration times, observations are also expedited by elimination of the need to increment the LO to image the entire band. This means simultaneous detection of spectral lines, simplified spectral index measurements, etc.

5 Technology description

The TKIP, Fig. 1a, consists of an electrically long section of superconducting transmission line, in coplanar waveguide (CPW) format. An intense “pump” waveform at angular frequency ω_P , and a weak “signal” waveform at frequency ω_S are guided into one end of the transmission line. Power is transferred from the pump waveform to the signal waveform as the two travel along the line. The signal emerges amplified at the other end of the line, and the pump is correspondingly attenuated.

The TKIP exploits the natural nonlinear behaviour of the surface inductance of the superconductor. The surface inductance (“kinetic inductance”) is due to the kinetic energy of the supercurrent. The series inductance per unit length, \bar{L} , of the transmission line depends on the current, I , as $\bar{L}(I) = \bar{L}(0)[1 + I/I_*^2]$, where I_* sets the scale of the effect.

The current on the transmission line due to the intense pump waveform makes the inductance of the line vary periodically. Thus, the non-linear inductance is the parameter that is used to generate the amplification (“parametric amplifier”). For the superconductors used,

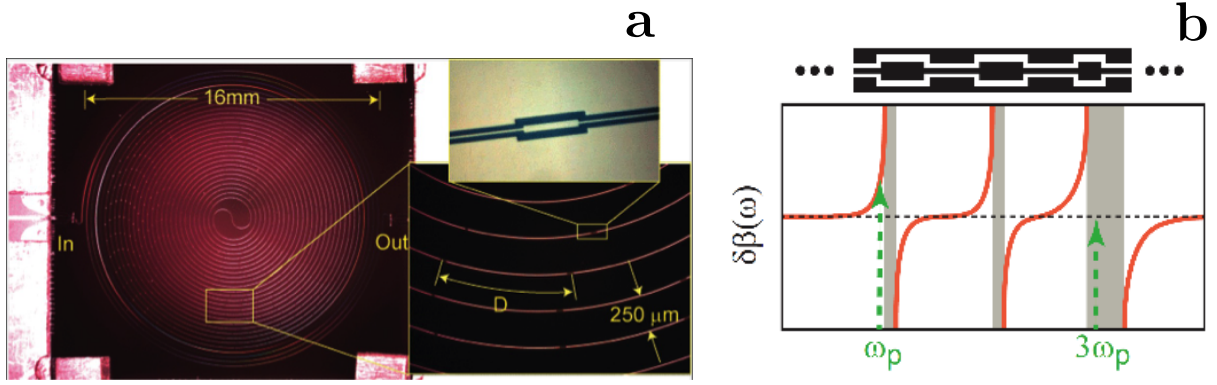


Figure 1: **a**, A picture of the microwave band TKIP, (μ TKIP), which consists of a 0.8 m length of NbTiN CPW line arranged in a double spiral to reduce resonances due to coupling between adjacent lines. The line is periodically loaded by widening a short section after every length $D=877$ micron as shown on the right, producing the stop band and dispersion characteristics. The phase velocity on the line is $0.1c$ due to its large kinetic inductance. **b**, An illustration of the effect of the periodic loading pattern (shown schematically) on the transmission of an infinite transmission line. The gray regions represent stop bands; waves in these frequency ranges decay evanescently. As the fractional width of the third stop band is much larger than the first, the pump can be placed at a propagating frequency while $3\omega_P$ is blocked.

the nonlinearity is almost purely reactive meaning that these devices are non-dissipative. Consequently, one, the TKIP should not dissipate power, and two, the TKIP need not produce thermal noise or shot noise and, in principle, the noise performance is limited only by quantum mechanics.

6 Theory of operation

The phase velocity depends on inductance, $v_p(I) = 1/\sqrt{\bar{L}(I)\bar{C}} \approx v_p(0)(1 - \alpha I^2/2I_*^2)$, where α is a measure of the kinetic inductance. This is analogous to the optical Kerr effect where the refractive index is intensity dependent. The intensity dependence of the refractive index results in the nonlinear process of four-wave mixing (FWM). Fiber optic paramaps, which operate at optical/near-IR frequencies, are based on this mixing [4] and coupled mode equations have been developed to quantify their parametric gain. We can use this theory [14] to predict the gain of the TKIP.

The FWM allows for two different pump waveforms, but we consider the degenerate case where the two are equal. When the pump and the signal are allowed to mix, an “idler” waveform is generated. The relationship between the frequencies of the three waveforms is by $\omega_S + \omega_I = 2\omega_P$ or $\omega_S + \omega_I = \omega_P$, where ω_I is the idler waveform angular frequency.

According this theory [14] the gain of a weak signal in the presence of a strong pump can

be written as,

$$G_S = 1 + \left[\frac{\gamma P_P}{g} \sinh g\ell \right]^2. \quad (1)$$

Here ℓ is the length of the transmission line, P_P is the pump power, and γ is a measure of the nonlinearity such that $\Delta\theta = \gamma P_P \ell$ is the power-dependent phase shift, in radians, resulting from the intensity dependent index. The parameter g depends on dispersion $\Delta\beta$,

$$g^2 = -\Delta\beta \left[\frac{\Delta\beta}{4} + \gamma P_P \right]. \quad (2)$$

and the dispersion in turn depends on the values of the propagation constant $\beta(\omega)$ at the signal, idler and pump frequencies through $\Delta\beta = \beta(\omega_S) + \beta(\omega_I) - 2\beta(\omega_P)$.

To a good approximation, a superconducting CPW transmission line is dispersionless, $\Delta\beta = \gamma = 0$. Thus the parametric gain in the signal is a quadratic function of the non-linear phase shift, $G_S = 1 + (\Delta\theta)^2$, since $\lim_{g \rightarrow 0} \frac{\gamma P_P}{g} \sinh g\ell = \gamma P_P \ell = \Delta\theta$.

For a dispersive transmission line, $\Delta\beta, g \neq 0$ and the presence of the $\sinh g\ell$ term in the gain expression can result in exponential gain. Maximum gain occurs when $\Delta\beta = -2\gamma P_P$ and ℓ is long enough so that $\Delta\theta \gg 1$. The gain is then $G_S \approx e^{2\gamma P_P \ell}/4 = e^{2\Delta\theta}/4$. Fiber optic paramps are able to access the exponential gain regime, producing gain > 60 dB, because they have builtin dispersion.

Dispersionless nonlinear media, such as our superconducting transmission line, support efficient harmonic generation. The formation of pump harmonics would severely limit the gain of the TKIP [7]. We therefore take the approach of introducing dispersion into the TKIP by periodically loading the line, Fig. 1b. The spacing between the loads is chosen to place a stop band around $3\omega_P$, the frequency of the lowest pump harmonic. Additionally, the loads are slightly offset to introduce dispersion features around the pump frequency itself, ω_P . This allows us to access the exponential gain regime. Thus, engineered dispersion simultaneously solves the problem of harmonic generation and provides access to the exponential gain regime.

7 Previous work

Non-linear response in electronic circuits has long been used to obtain parametric amplification. In the 1950s, for example, low-loss varactor diodes with voltage-variable capacitance allowed the construction of amplifiers operating in the microwave to millimeter-wave bands [10]. The nonlinear inductance of Josephson junctions in superconducting circuits can represent a nearly ideal nonlinear reactance, which led to strong interest in the development of Josephson-junction amplifiers [2, 12], and by the late 1980s improved junction fabrication techniques allowed sophisticated devices to be produced reliably, and narrow-band microwave amplifiers with excellent noise performance approaching the QL were demonstrated [9, 13, 19].

A drawback of early implementations of Josephson paramps was that the amplification was only achieved over a narrow bandwidth. Tuned circuits used to enhance the effect of

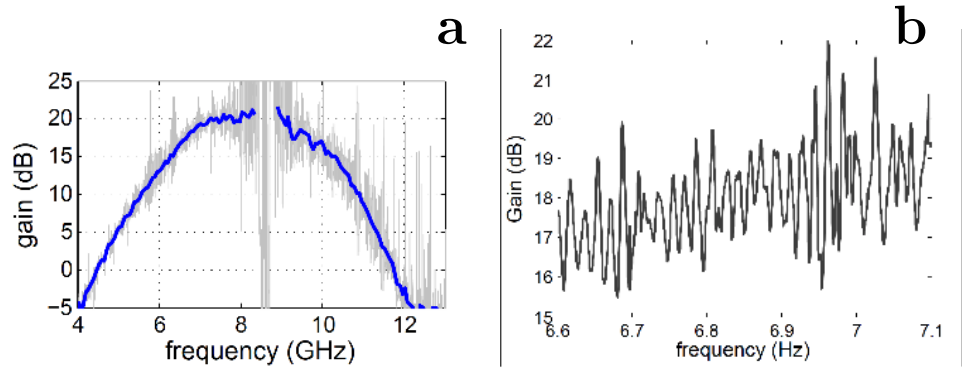


Figure 2: **a**, Measured gain of a μ TKIP using a 0.8 m NbTiN CPW (gray line). The measured gain is smoothed over 40 MHz to produce the blue line. **b**, Detail of gain curve between 6.6 and 7.1 GHz. The rapid gain variations are due to reflections at the ends of the device. We expect to reduce these reflections through better impedance matching and by implementing ground ties to connect the CPW grounds.

nonlinearity produced intrinsically narrow-band devices, unsuitable for applications needing wide instantaneous bandwidth. An alternative design [1, 15–17] dispensed with the resonant circuit and the device took the form of a nonlinear transmission line with either a distributed nonlinearity or a series of discrete nonlinear elements. Though intrinsically wide band, it did not result in practical devices because of dispersive effects and problems associated with producing the required degree of nonlinearity.

The JPL group constructed a TKIP with significant gain in the microwave frequency band 8 - 14 GHz (μ TKIP). The success of that device was due to the use of NbTiN superconductor,

which exhibits strong nonlinearity and low loss. The group reported its breakthrough in the journal *Nature Physics* [5]. The publication makes clear what improvements need to be made in order to realize the full potential on this new technology. They include, one, increasing the gain and eliminating a fine-scale variation of the gain; two, understanding and suppressing noise in excess of the QL. The group has made progress toward these improvements and the project I propose will extend their work. Presently, the group has designed and fabricated a first iteration of a W-band TKIP (wTKIP) using, in part, seed funds from an NRAO ALMA technology development grant, but these funds were exhausted before testing could commence.

8 Approach

The project will be split up into two task groups. One is to construct a test set up for the wTKIP and the other is to develop the amplifier. Amplifier development consists several sub-tasks.

8.1 Construction of Test Setup for W-Band TKIP

I propose to construct a test setup capable of measuring the gain and noise of the prototype W-band TKIP (wTKIP). The noise measurement would be done via the Y-factor method, which uses a variable temperature termination. This will involve installing waveguides in a cryostat with a 1 K cryocooler and a cryogenic HEMT amplifier. To avoid saturation, the ability to null the pump tone before it enters the HEMT, will be built into the system. An existing lab down converter system will be used to shift the output band to a range that can be measured by a microwave spectrum analyzer.

8.2 Amplifier development

Amplifier development will begin by testing the first iteration wTKIP. Improved coupling to the wTKIP and the use of ground plan straps will be explored to improve the amplifiers's gain characteristics. Excess thermal noise in the wTKIP will be suppressed through better heat sinking.

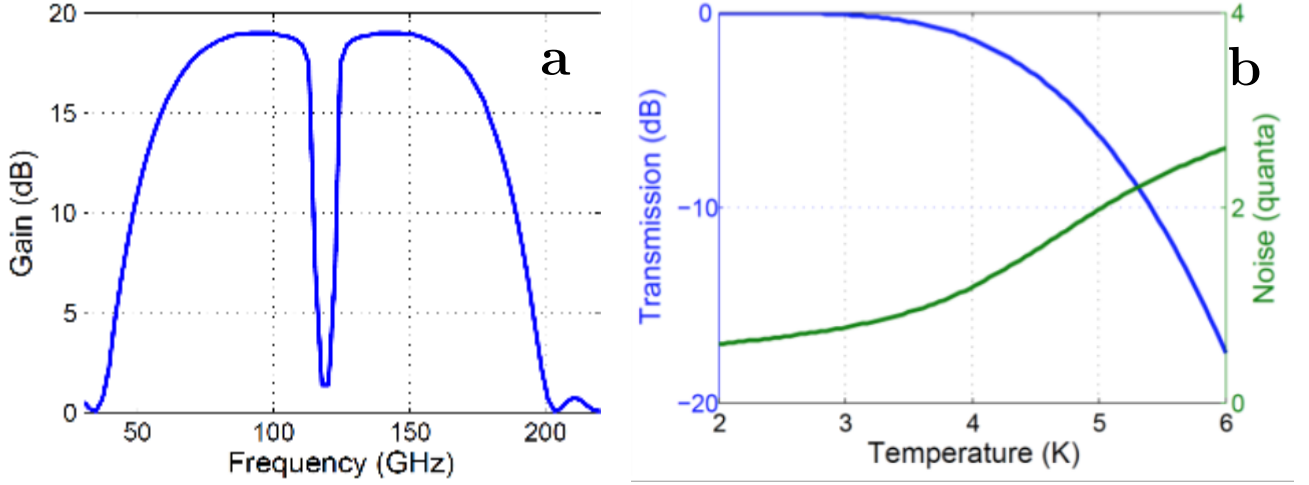


Figure 3: **a**, Calculated gain of the W-band TKIP (wTKIP), assuming a pump power half as large as what was used with the microwave TKIP described in [5] (μ TKIP). The dip in the center is due to dispersion around the pump frequency. **b**, Calculated transmission through a 100λ section of the NbTiN line versus temperature in blue. The green line shows the estimated noise contribution of the device, assuming no non-thermal noise sources.

8.2.1 Test wTKIP

The wTKIP uses the same CPW geometry as the μ TKIP [5], but the periodic structure has been adjusted for a pump at 118 GHz. The overall length of the transmission line is decreased, but the length measured in wavelengths is somewhat increased as the operating frequency is ten times higher. Fig. 3a shows the calculated gain based on the coupled mode theory, assuming a pump power, P_P , half of that used for the microwave device.

We will determine the sensitivity (noise), gain characteristics of the first generation wTKIPs and make necessary design corrections to enhance their performance.

8.2.2 Ground plane straps

A typical gain curve is shown in Fig. 2. While the average gain displays a broad maximum, reaching 20 dB, there is a fine-scale variation of the gain due to standing waves created by reflections at the ends of the line. These ripples were more severe in earlier μ TKIP and have been reduced through better impedance matching. A further improvement would be to connect the CPW ground planes with “straps” to prevent the excitation of a slot line mode on the CPW. Such straps could be implemented by bonding superconducting connections between the ground planes after fabrication, or by fabricating superconducting air bridges.

This project explore the use of straps to suppress the gain ripple.

8.2.3 Improvement of coupling to TKIP

For the first iteration wTKIP, we will use commercial waveguide to WR10 transitions and WR10 sparkplug connectors to bring the millimeter-wave power into and out of the device housing. The power is brought onto the chip using wire bonds. While these devices will be sufficient for initial testing, reflections are likely to be a limitation. Further iterations of the wTKIP will include on-chip waveguide probes and reflectionless filters.

Waveguide probes: To further reduce gain ripple and improve performance, it will be necessary to improve signal coupling into and out of the paramp. Full band waveguide probes coupled to transmission lines have been developed for use up to THz frequencies. An example of a simple design is the radial probe [18] which has a feed impedance of approximately 20 Ohms. The Microdevices Lab at JPL has previously manufactured similar probes for SIS mixer based instruments (e.g. [6]). The impedance of the TKIP CPW transmission line is typically 50 Ohms or larger so coupling to these probes requires an impedance matching section which potentially limits the bandwidth. As used in the MUSIC [3] array, another option for the TKIP is to couple incoming radiation through a set of multiple long slot antennas which are combined in a phased array. This potentially gives wide bandwidth and the feed impedance of the antennas is closer to the impedance of the TKIP transmission line. I propose to design and test both types of optical couplings as part of the postdoctoral appointment.

Reflectionless filter: The very wide bandwidth of the TKIP could be a problem in a waveguide system where, for example, frequencies below the waveguide cutoff are reflected. While the gain of the TKIP is directional, there is no isolation between output and input as reverse directed power is almost perfectly transmitted back through the transmission line. An additional reflection at the TKIP input would give rise to an instability if the input and output reflection coefficients, Γ_{in} , Γ_{out} , satisfy the condition $\Gamma_{in}\Gamma_{out}G_S > 1$.

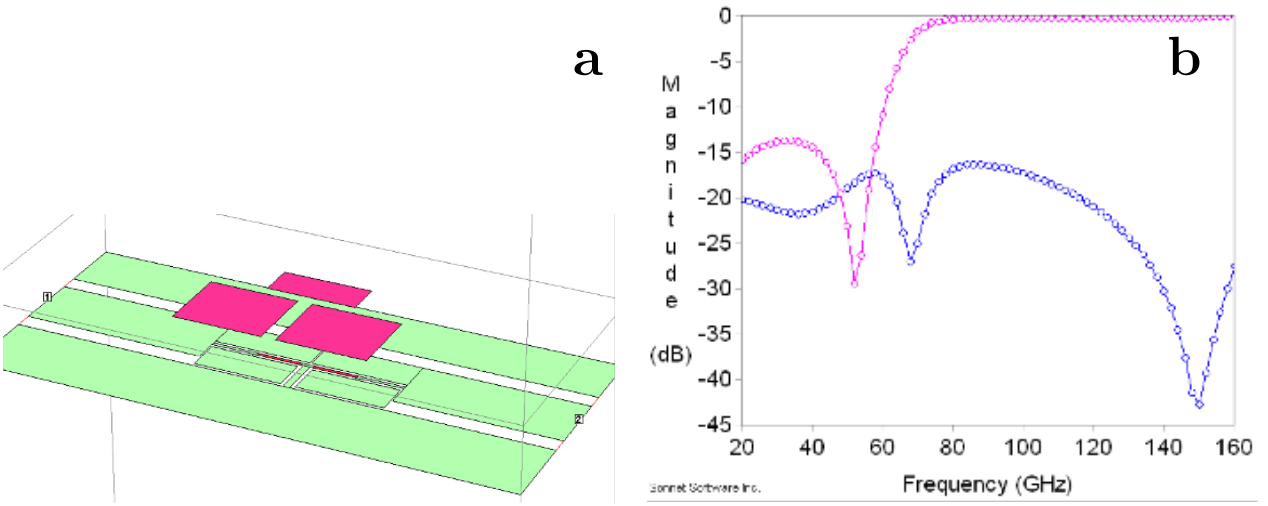


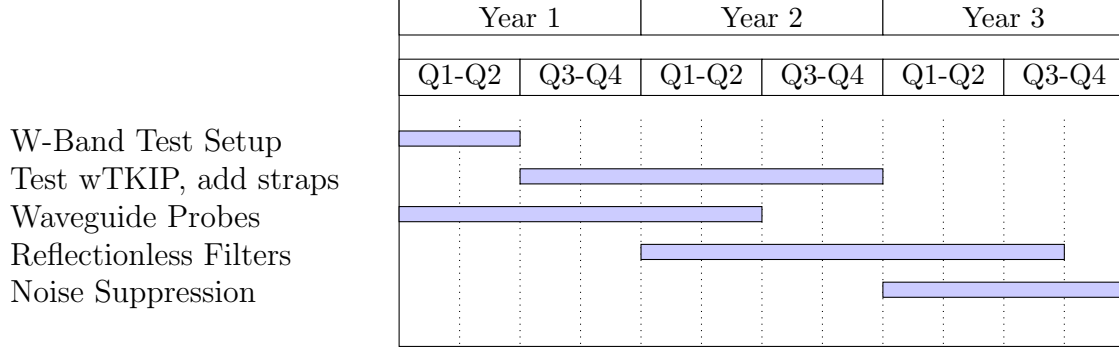
Figure 4: **a**, Geometry of a reflectionless 60 GHz high pass filter. Layers: Green = NbTiN, red = resistor, magenta = metal conductor. An SiO₂ layer isolates the NbTiN and conductor layers. The design is low risk in that no contact is required between NbTiN and the top conductor and that the smallest critical dimension required is 2 microns. The thickness of the NbTiN is the same as that of the paramp. **b**, Simulated reflection S11 (blue) and transmission S21 (red) of the filter.

To deal with this potential problem, we propose to develop millimeter-wave reflectionless filters to dissipate the out of band power. Reflectionless filters have been implemented at lower frequencies [8] and are a simpler solution than a diplexer with one terminated output, where the control of input match over the entire band may not be the first consideration. Simulations of this approach suggest its feasibility. Our design uses four fabrication layers. The NbTiN base layer defining the CPW structure would be the same as the paramp CPW layer so that the filter and the paramp may be fabricated on the same chip Fig. 4. The maximum simulated reflection of the filter between 20 and 160 GHz is -16.5 dB, which would correspond to instability at 33 dB gain and above. I propose to optimize this filter design as part of the postdoctoral appointment and then fab and test a prototype.

8.2.4 Suppression of excess noise

Since their publication [5], the JPL group determined that the excess noise in their μ TKIP, approximately 3 or 4 photons in excess to the QL of $1/2$ photon (per second per Hz bandwidth), was thermal in origin. They came to this conclusion by measuring the noise of the μ TKIP while pulsing the pump waveform with on-time ranging from 0 to 100 percent. They found that the excess noise decreases as the on-time decreases. Further, they found that the

temperature of the μ TKIP chip increases with pump power. These observations indicate that the amplifier is not operating in a completely dissipationless mode and that at least a component of the observed excess noise is thermal noise. At 100 GHz, we expect that the wTKIP can still be essentially quantum limited at cryocooler temperatures Fig. 3b, but the chip must be prevented from significantly heating. A goal of this project will be to study the effect of better heat sinking on the excess noise of the wTKIP.



9 Work plan

The Gantt chart outlines the scheduling of the proposed work. Completion of the W-band test setup is a milestone, enabling the test and characterization of the wTKIP as well as the testing of coupling match solutions into and out of the paramp. Another key milestone is conclusion of the wTKIP testing, at which time we will be in a position to propose a more advanced paramp design that integrates our coupling reflectionless filter ideas on chip. After year two, I will focus attention on achieving quantum limited noise at high P_P and at higher operating temperature, 2 K, for easy compatibility with pulse tube coolers. I expect to publish the performance of the early iteration wTKIP in year two and then publish the results of the coupling and noise improvements in year three.

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