Energy Balance and Climate

EES 3310/5310

Global Climate Change

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Class #3: Monday, January 24 2022

Announcement

For lab (before you go to lab today, or when you get to lab)

- Go to the Lab #1 assignment on the course web site.
  - Click on the link to accept the assignment at GitHub Classroom:

Assignment

Accept the assignment at GitHub Classroom at https://classroom.github.com/a/5vrnN29K.

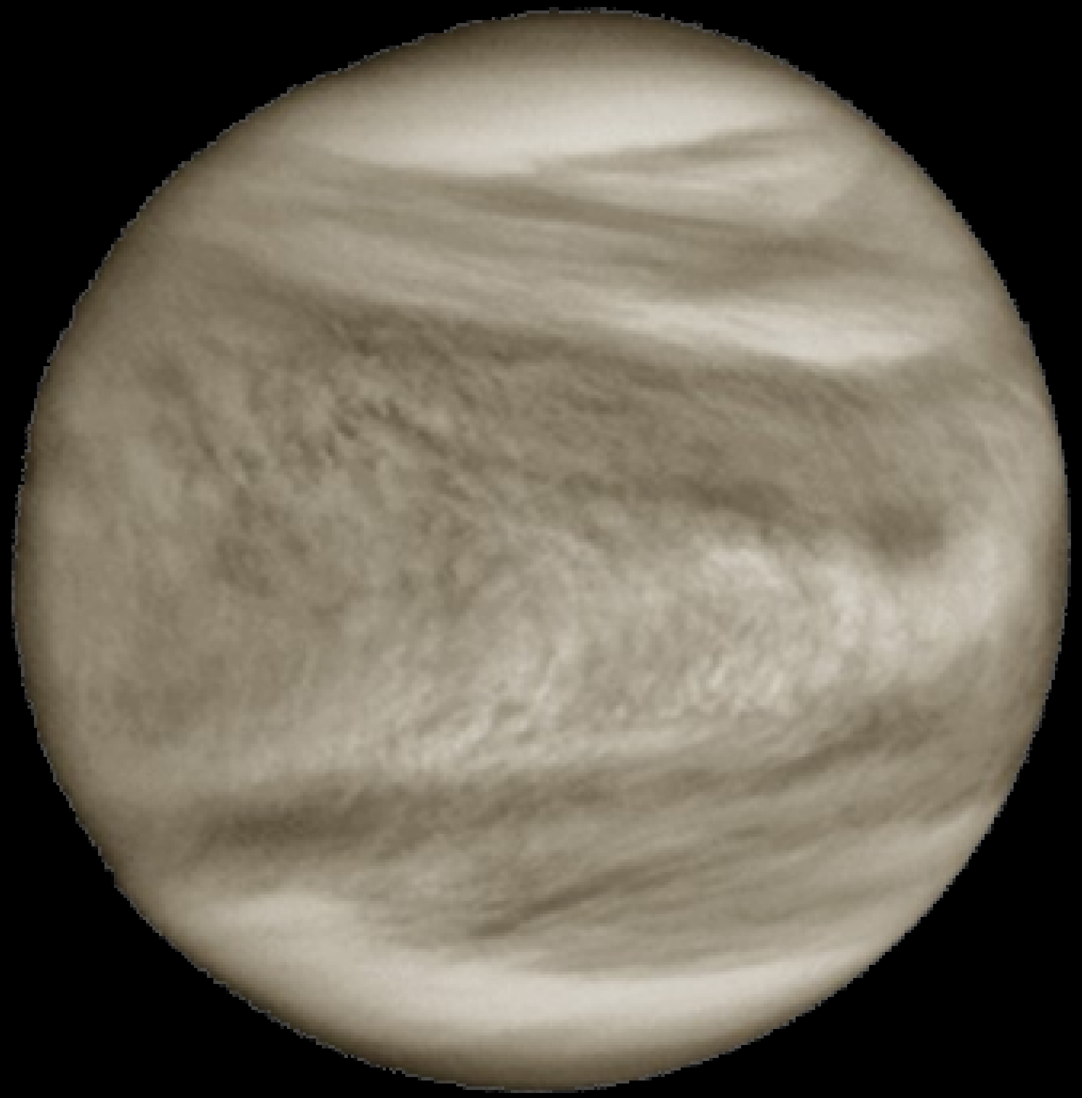
Looking for a Good Home

Bad

-28°F

Good

71°F



Worst

800°F

Basic Concepts

Vocabulary

- **Energy, Heat:**
  - Heat = energy flowing spontaneously from hot to cold
- **Power:** speed at which energy flows or transforms

$$\text{Power, Flux} = \text{Heat flow}/\text{Time}$$

$$\text{Heat, Energy} = \text{Power} \times \text{Time}$$

- **Intensity:** Concentration of power

$$\text{Intensity} = \text{Power}/\text{Area}$$

$$\text{Power} = \text{Intensity} \times \text{Area}$$

Temperature of a planet

- **Basic principle:**

  Steady temperature if and only if

  $$\text{Heat Power}_{\text{in}} = \text{Heat Power}_{\text{out}}$$

- **How can heat get in or out?**
  - Electromagnetic radiation

Electromagnetic Waves

- **Color** and **brightness**
  - **Color:**
    - Two ways to measure **color**:
    - Wavelength (λ)
    - Wavenumber ($n = 1/\lambda$)
  - Archer mostly uses wavenumber
    - Math is simpler that way
  - **Brightness:**
    - Intensity (power/area, Watts/square meter)

Colors

Color	wavelengths	wavenumbers
infrared	> 0.70 μm	< 14,000
red	~ 0.70–0.64 μm	~ 14,000-16,000
orange	~ 0.64–0.59 μm	~ 16,000-17,000
yellow	~ 0.59–0.56 μm	~ 17,000-18,000
green	~ 0.56–0.49 μm	~ 18,000-20,000
blue	~ 0.49–0.45 μm	~20,000-22,000
violet	~ 0.45–0.40 μm	~22,000-25,000
ultraviolet	< 0.40 μm	> 25,000

All you need to think about is
shortwave vs. **longwave** radiation.

Shortwave and **longwave**:

- **Shortwave**:
  - Near-infrared, visible, ultraviolet
  - $\lambda < 3\mu\mathrm{m}$
  - $n > 3,300\mathrm{cm}^{-1}$ (cycles per centimeter)
- **Longwave**:
  - Mid-infrared, far-infrared
  - $\lambda > 3\mu\mathrm{m}$
  - $n < 3,300\mathrm{cm}^{-1}$

More on this on Wednesday ...

Light and the Atmosphere

- **Shortwave**
  - Very hot things
  - Sun is almost 10,000° F
  - Light bulbs are about 5,000° F
- **Longwave**
  - Cooler things
  - Earth is about 68° F on average

Photo credit: Burak K, Pexels

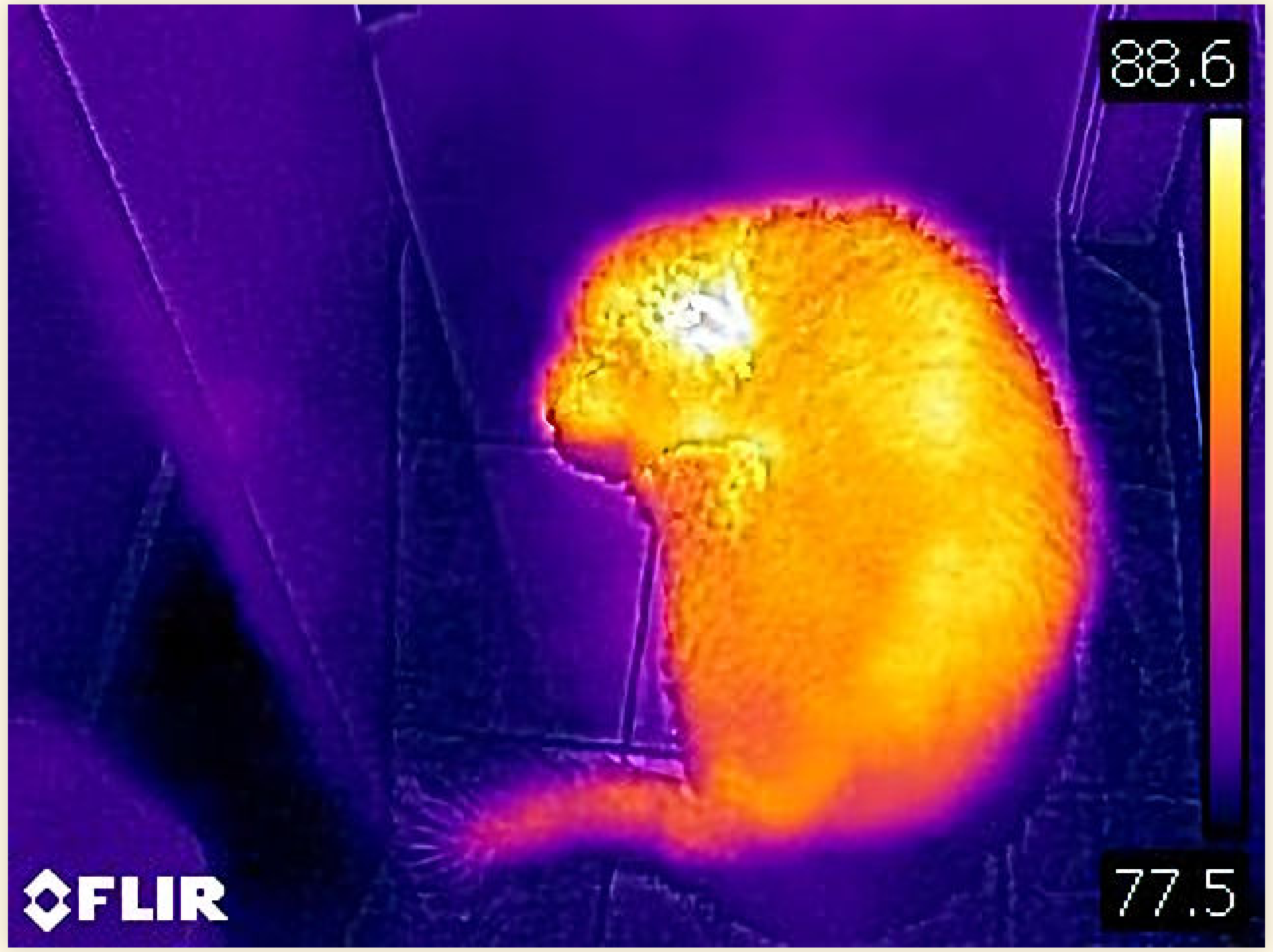
Photo credit: Kateryna Babaieva on Pexels

Photo credit: Yongrow Medical

4 Laws of Radiation

1. All objects continually radiate energy
2. Hotter objects are brighter
3. Hotter objects radiate at shorter wavelengths
4. Objects that are good absorbers are also good emitters
    - Black objects emit & absorb the most
    - Transparent and white objects emit & absorb the least

Example of Radiant Heat



- Featuring my dog, Finley.
- All objects emit electromagnetic radiation
- Hotter objects emit
  - More intense the radiation
  - Shorter wavelengths

Blackbody Radiation

Blackbody Radiation

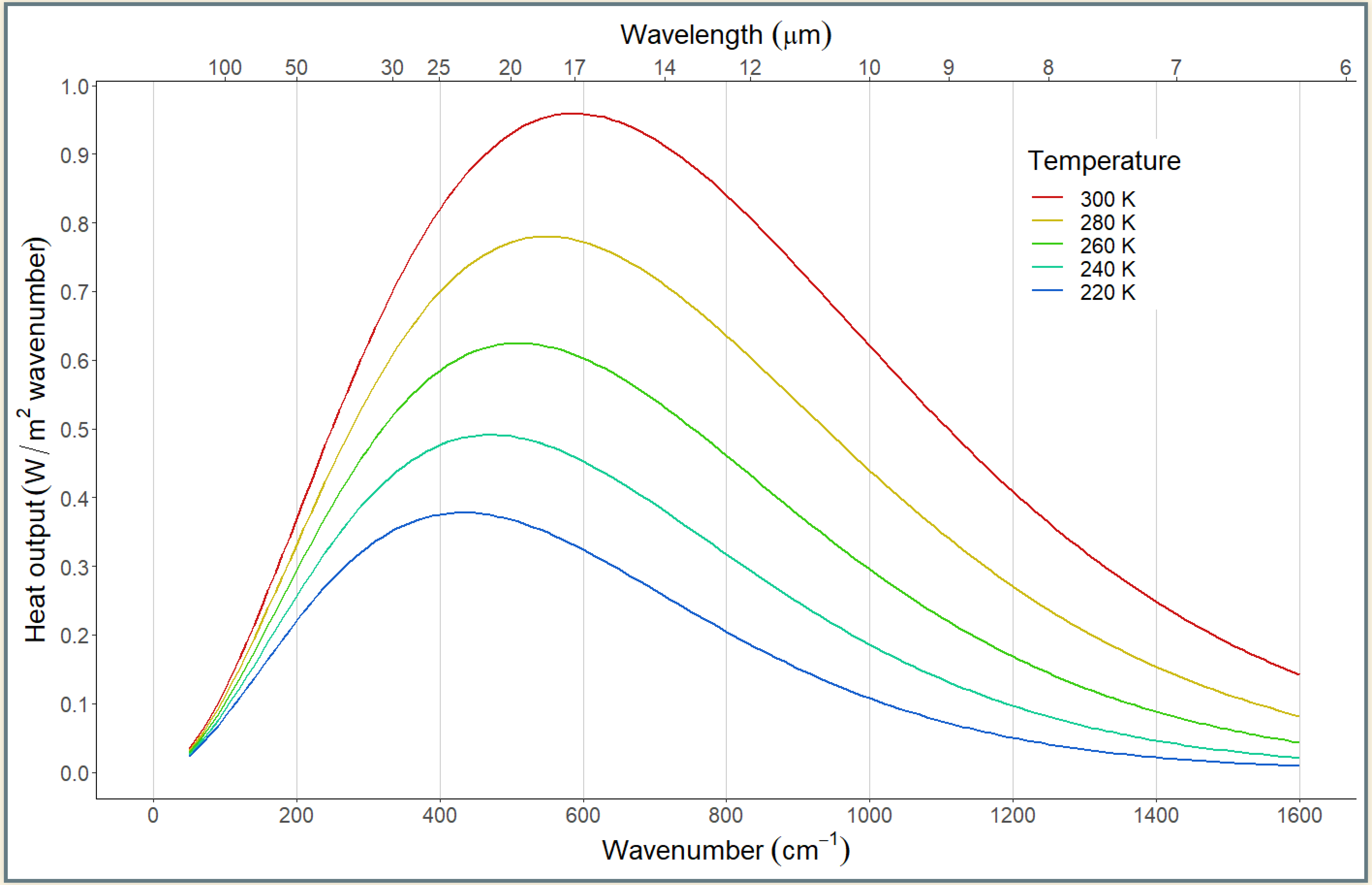
Emissivity (ε) measures how black something is:

- $\varepsilon = 1$ for perfectly black
- $\varepsilon = 0$ for perfectly white or transparent
- In between for gray.
- Black, white, and gray: ε is the same for all wavelengths.
- Colored objects: ε is different for different wavelengths.
- **For simplicity:** start by assuming everything is black, white, or gray.

Remember: Good emitters are good absorbers

Fundamental rule: Temperature and emissivity determine radiation.

Heating Up: What Changes??



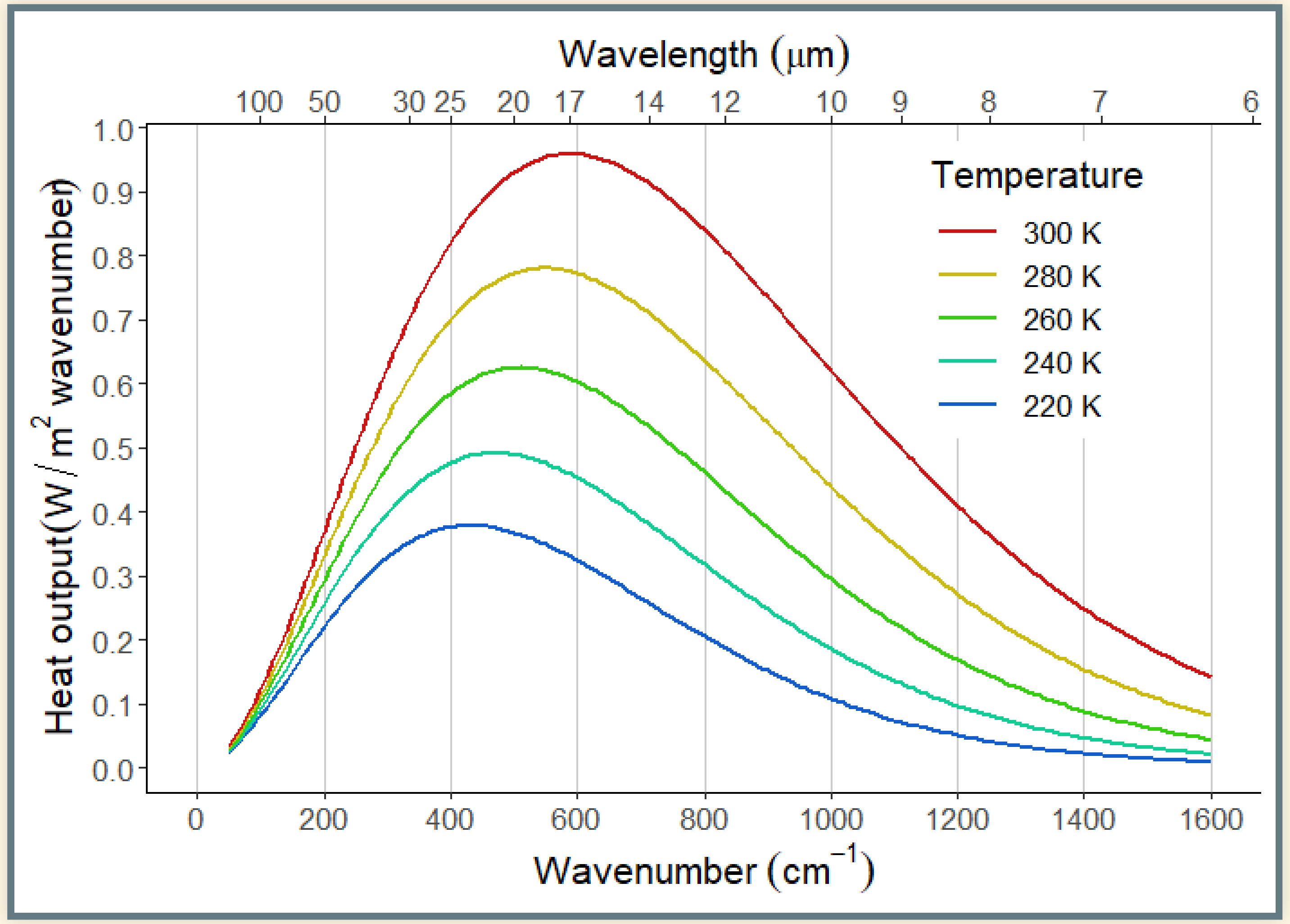
Heating Up: What Changes?

- Hotter temperature:
  - Brighter (greater intensity)
  - Bluer (greater wavenumber, shorter wavelength)

A curious thing:

- A hot black object glows with color!

Total intensity =
area under curve



Mathematical Description

Blackbody Radiation

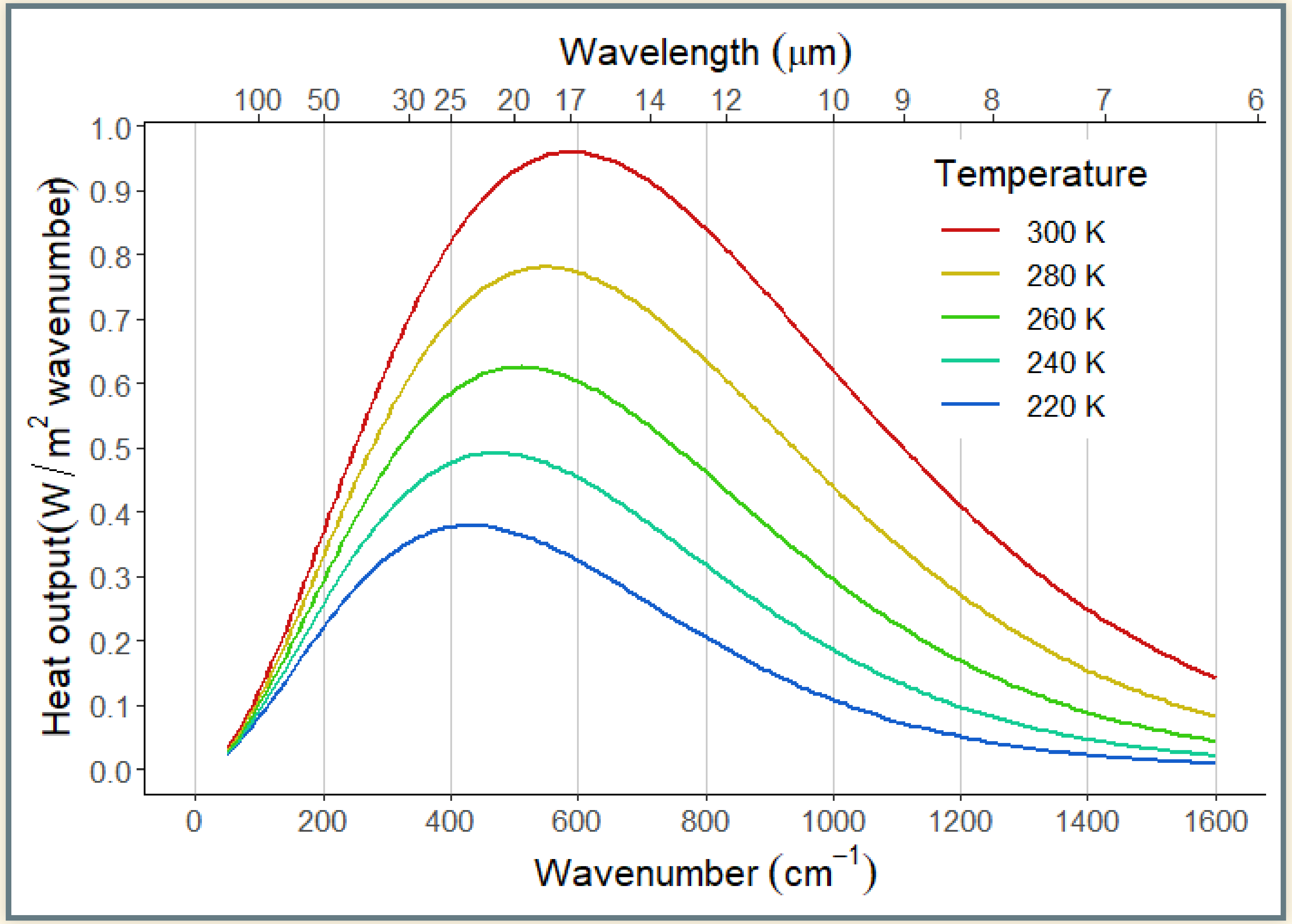
Intensity (brightness):

Stefan-Boltzmann law

$$I = \varepsilon \sigma T^4$$

after Josef Stefan
and Ludwig Boltzmann

- ε = emissivity
  - Different for different objects.
- σ = Stefan-Boltzmann constant.
- T = absolute (Kelvin) temperature.

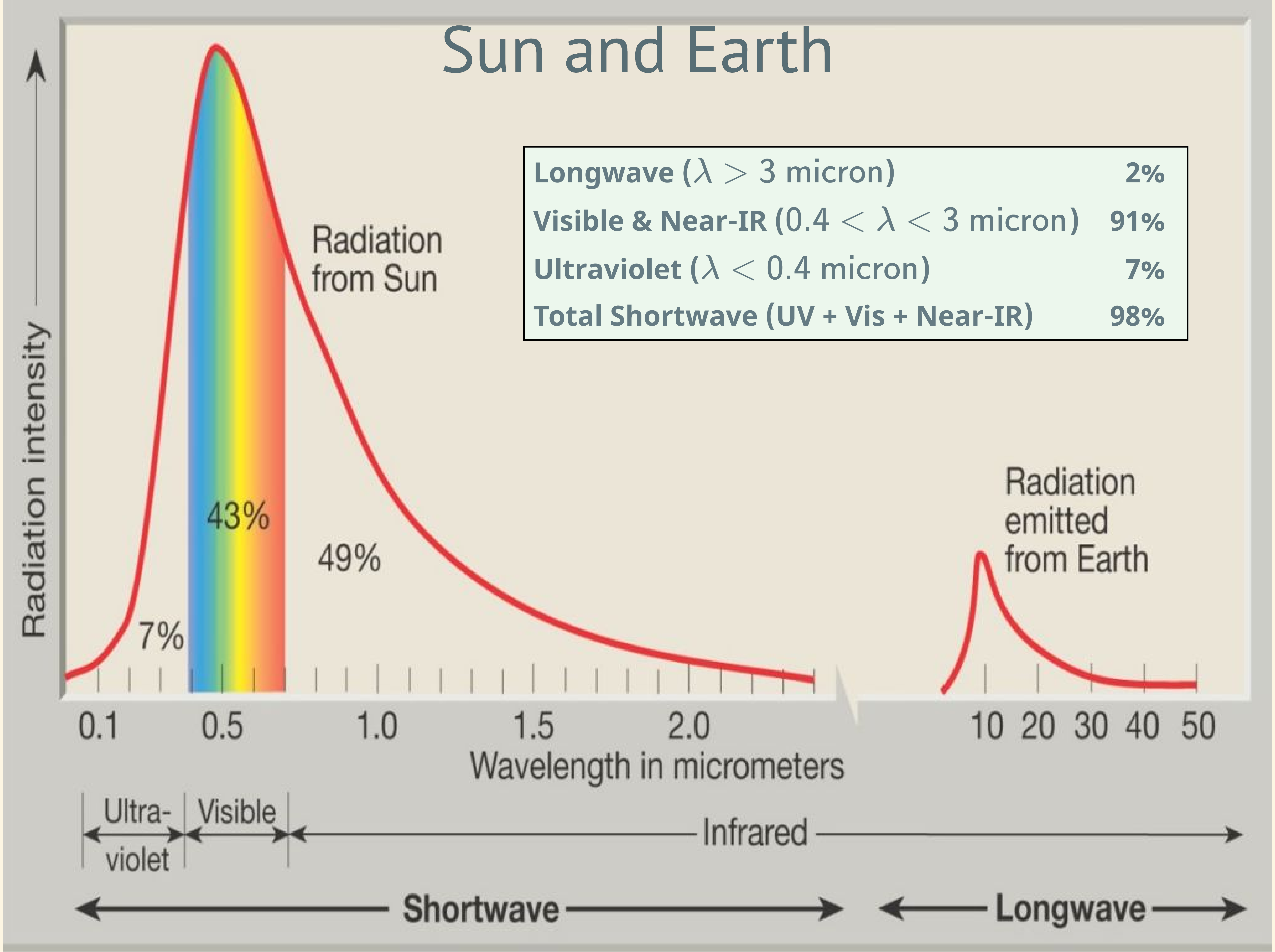


Color: Peak wavenumber proportional to (Kelvin) temperature.

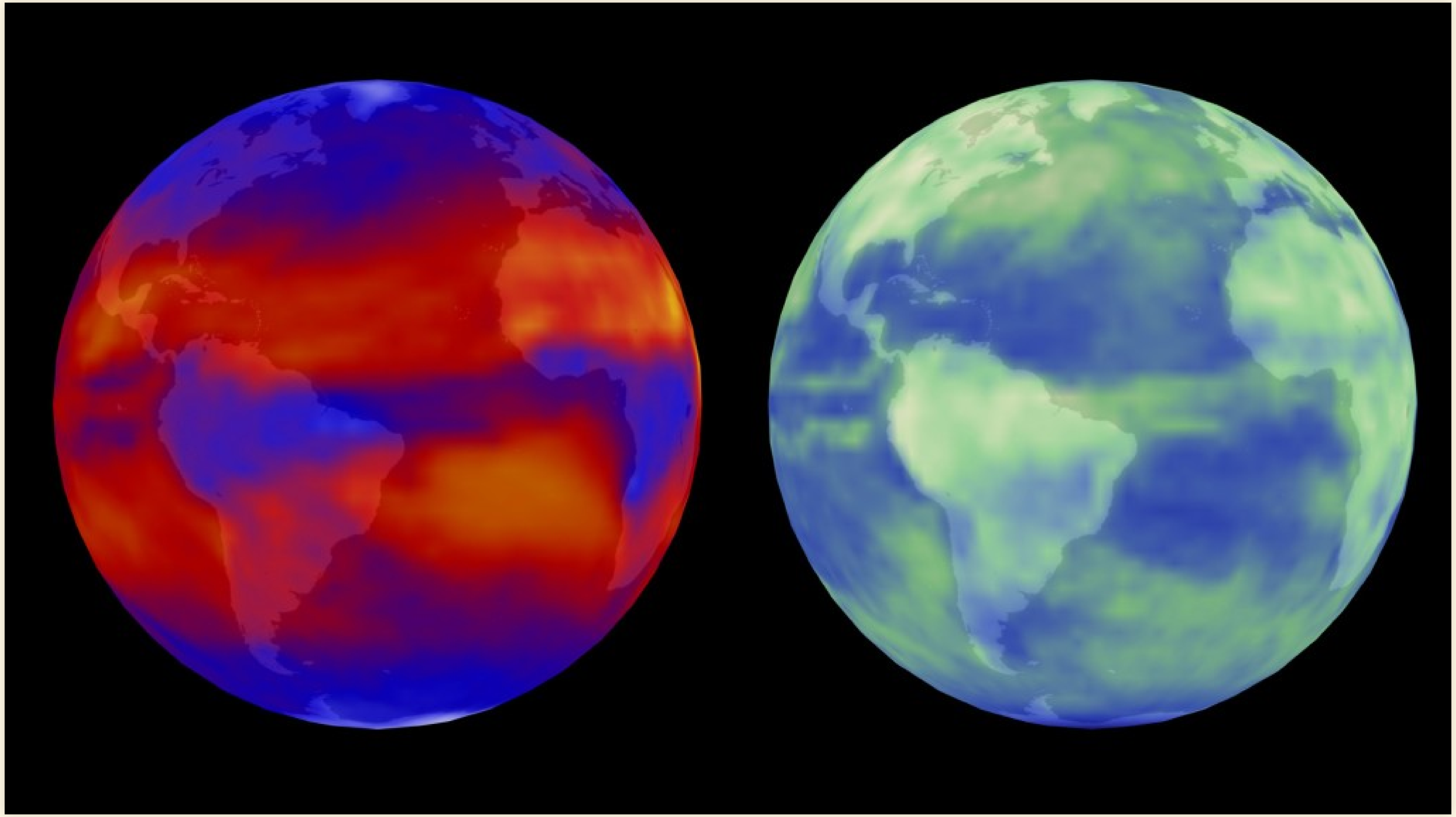
Helpful Hint:
Fourth power on a calculator: press the x^2 button twice.

Sun and Earth

Longwave ($\lambda > 3$ micron)	**2%**
Visible & Near-IR ($0.4 < \lambda < 3$ micron)	**91%**
Ultraviolet ($\lambda < 0.4$ micron)	**7%**
Total Shortwave (UV + Vis + Near-IR)	**98%**



Earth and Radiation



False-color images of radiation from Earth, seen by NASA Terra satellite:

- Left: Thermal radiation (blue → red → yellow = dim → bright)
- Right: Reflected sunlight (blue → green → white = dim → bright)

Efficiency of Light Bulbs

Type of Bulb	Efficiency
Standard 40W	1.8%
Standard 60W	2.1%
Standard 100W	2.6%
Quartz Halogen	3.5%
Ideal black body @ 7000K	14.0%
Compact Fluorescent	8–12%
LED	20–44%

- 7000K is the optimal temperature for a black body to emit visible light, but it will melt every known substance.
- Standard light bulbs operate at around 2000–3300 K.

Calculating Earth's Temperature: Bare-Rock Model

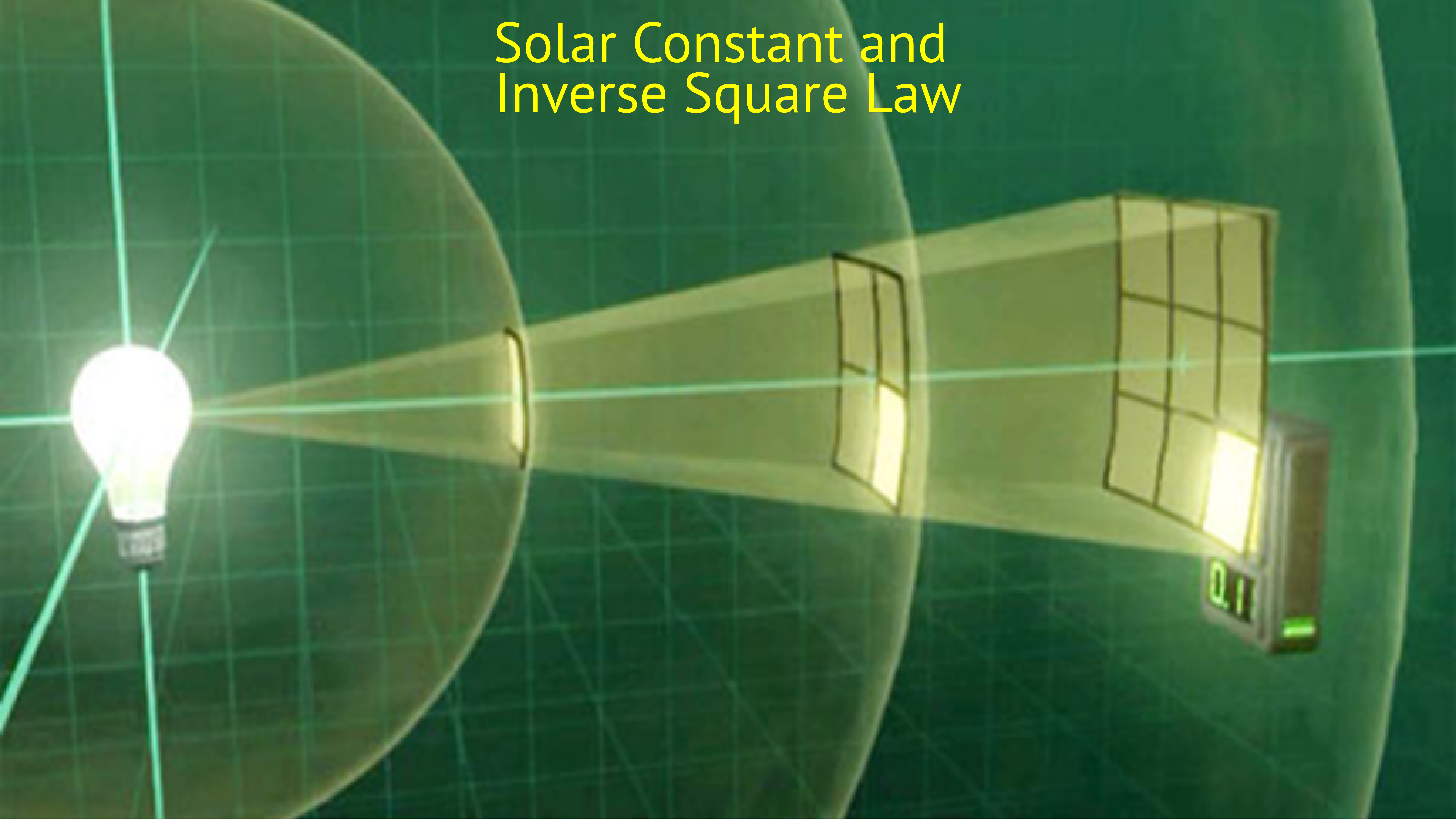
Basics

Steady Temperature

- Heat in must balance heat out
- $\text{Total Heat Flux (Power)} = \text{Area} \times \text{Intensity}$
  - Total heat flux in (F_{in}):
    - Intensity depends on solar constant and albedo
    - Does not depend on earth's temperature
  - Total heat flux out (F_{out}):
    - Intensity depends on earth's temperature and emissivity
- Strategy:

1. Figure out F_{in}.
2. Figure out temperature T that makes $F_{\text{out}} = F_{\text{in}}$.

Solar Constant and Inverse Square Law

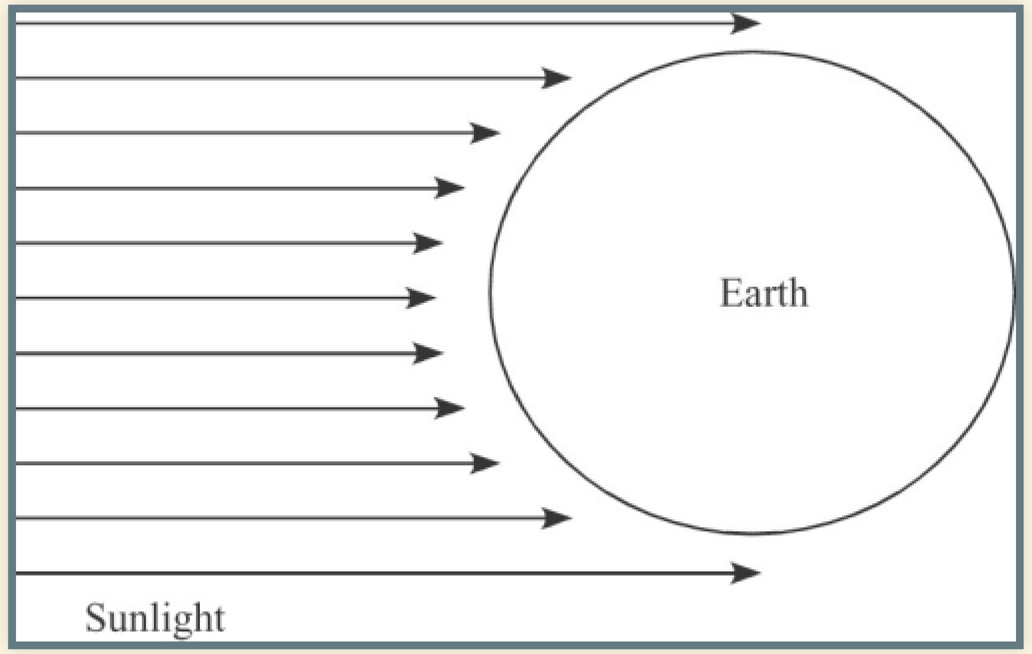


What is F_{in}?

- $F_{\text{in}} = \text{Area} \times \text{Intensity absorbed}$
  - Intensity absorbed $= (1 - \alpha) \times I_{\text{in}}$
    - $I_{\text{in}} = 1350 \text{ W/m}^2$
    - Average albedo $\alpha = 0.30$ (30% of sunlight is reflected)

What is area?

- Area = silhouette or shadow
- Circle: πr^2



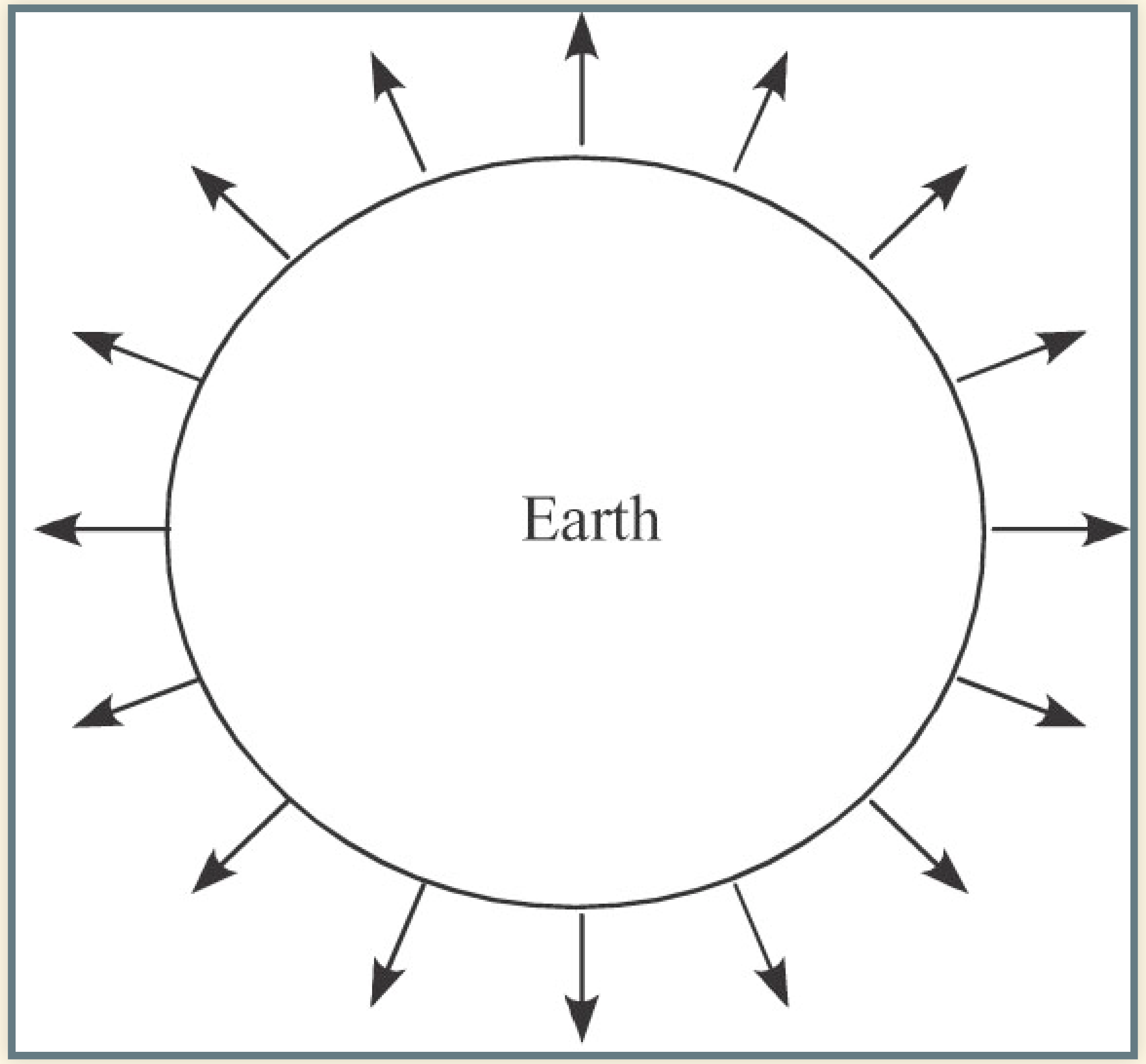
What is F_{in}?

- $F_{\text{in}} = \pi r_{\text{Earth}}^2 \times (1-\alpha) I_{\text{in}}$
  - $\pi r^2 = 1.3 \times 10^{14} \text{m}^2$
  - $\alpha = 0.30$
    - $(1-\alpha) = 0.70$
  - $I_{\text{in}} = 1350 \text{ W}/\text{m}^2$
- $F_{\text{in}} = 1.3 \times 10^{14} \text{ m}^2 \times 0.70 \times 1350 \text{ W}/\text{m}^2$

  $= 1.2 \times 10^{17} \text{Watts}$
  - 11,000 times total human energy production.

What is F_{out}?

- $F_{\text{out}} = \text{Area} \times I_{\text{out}}$
  - $I_{\text{out}} = \varepsilon \sigma T^4$
    - $\varepsilon = 1$ (blackbody)
    - $\sigma = 5.67 \times 10^{-8}\ \text{W/m}^2/\text{K}^4$
  - What is area?
    - Sphere: $4\pi r^2$
  - $F_{\text{out}} = 4\pi r_{\text{earth}}^2 \times \varepsilon \sigma T^4$



Putting it all together

$$F_{\text{out}} = F_{\text{in}}$$

$$4\pi r^2 \times \varepsilon\sigma T^4 = \pi r^2 (1-\alpha) I_{\text{in}}$$

$$4\pi r^2 \times \varepsilon\sigma T^4 = \pi r^2 (1-\alpha) I_{\text{in}}$$

$$4\varepsilon\sigma T^4 = (1-\alpha) I_{\text{in}}$$

$$T^4 = \frac{(1-\alpha) I_{\text{in}}}{4\varepsilon\sigma}$$

- Total flux (power) radiated from sun doesn't change with distance.
- At a distance r total flux spreads over sphere of radius r
- Intensity = Total Flux / Area:
  - Proportional to $1/r^2$.
- At edge of Earth's atmosphere, solar intensity = $1350\ \mathrm{W/m^2}$.

$$T = \sqrt[4]{\frac{(1-\alpha) I_{\text{in}}}{4\varepsilon\sigma}}$$

Temperature of Earth

Temperature of Earth

- Steady Temperature:
  - Heat flux in must balance heat flux out ($F_{\text{out}} = F_{\text{in}}$).
  - F_{in}:
    - Does not depend on earth's temperature.
    - Depends on solar constant and earth's albedo.
  - F_{out}:
    - Depends on earth's temperatu
  - T adjusts until heat out = heat i

Helpful hint:

To take the fourth root on a calcula
press the square-root key ($\sqrt{}$) twic

$$T = \sqrt[4]{\frac{(1-\alpha)I_{\text{in}}}{4\varepsilon\sigma}}$$

Temperature of Earth

$$T = \sqrt[4]{\frac{(1-\alpha)I_{\text{in}}}{4\varepsilon\sigma}}$$

Earth:

(Note: My numbers are slightly different from Archer's textbook)

- $I_{\text{in}} = 1350\ \text{W}/\text{m}^2$
- $\alpha = 0.30$
- $\varepsilon = 1$
- $\sigma = 5.67 \times 10^{-8}\ \text{W}/(\text{m}^2\text{K}^4)$
- Calculate T:
- $T = 254\ \text{K} = -19^\circ\text{C} = -2^\circ\text{F}.$

Temperature of Earth

$$T = \sqrt[4]{\frac{(1-\alpha)I_{\text{in}}}{4\varepsilon\sigma}}$$

Earth:

(Note: My numbers are slightly different from Archer's textbook)

- $I_{\text{in}} = 1350\ \text{W}/\text{m}^2$
- $\alpha = 0.30$
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- $\sigma = 5.67 \times 10^{-8}\ \text{W}/(\text{m}^2\text{K}^4)$
- $T = 254\ \text{K} = -19^\circ\text{C} = -2^\circ\text{F}.$

How does this compare to Earth's actual temperature?

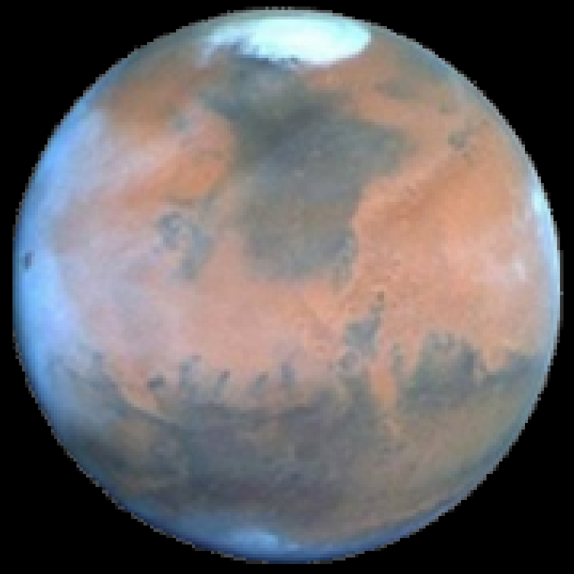
Comparing Theory and Observation

Radiative Temperature

- Satellites orbiting in space can measure longwave radiation from earth
- To the satellites, the earth looks very much like a blackbody at the bare-rock temperature (254 K).
- Thus, scientists generally call the bare-rock temperature the **radiative temperature** because it describes the radiation coming off the earth.
- However, the surface temperature of the earth is around $295\ \mathrm{K} = 71^{\circ}\mathrm{F}$, which is significantly different from the radiative, or bare-rock, temperature.

Terrestrial Planets

The Terrestrial Planets

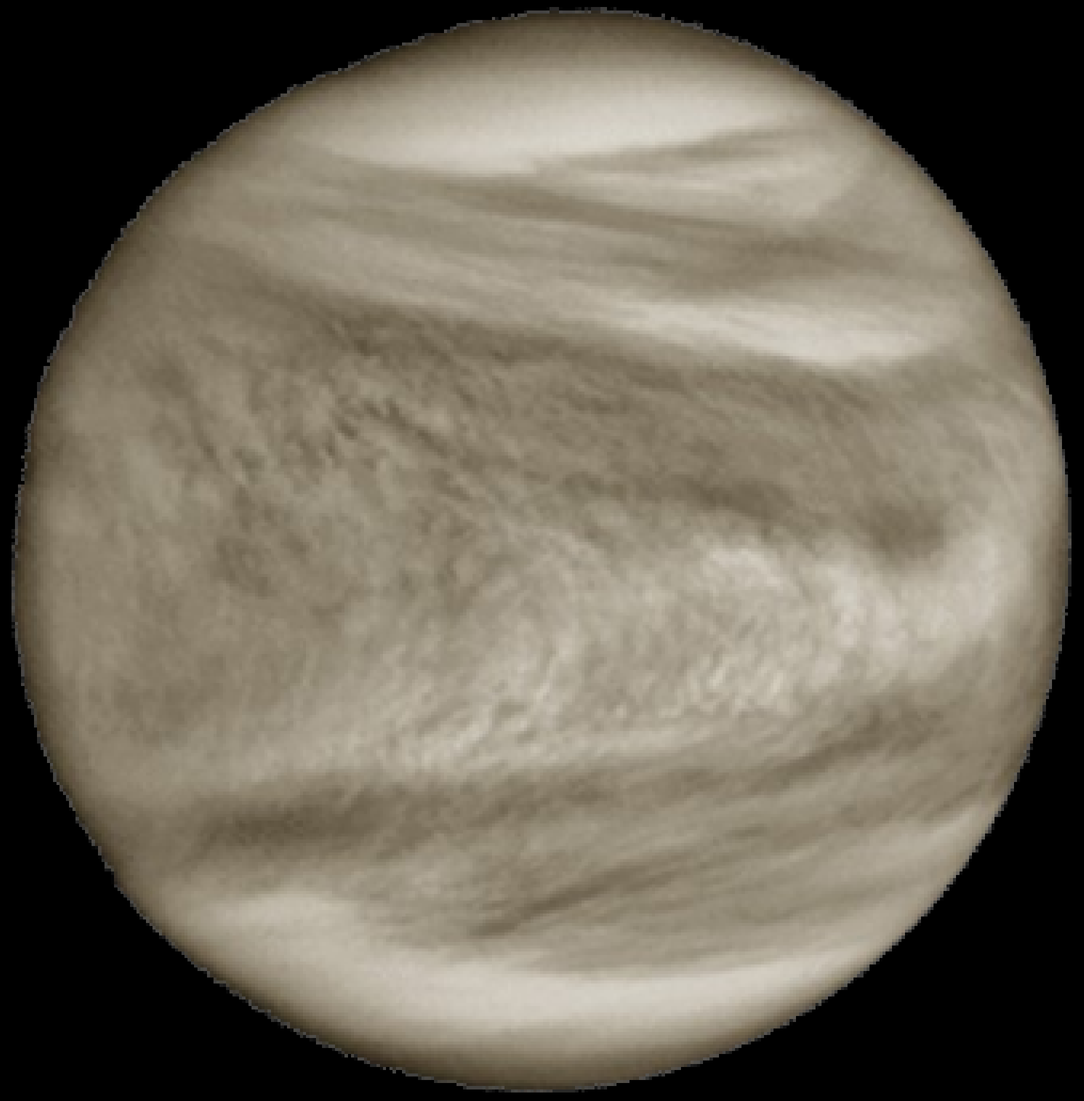


Mars

240 K

Earth

295 K



Venus

700 K

Terrestrial Planets

	Earth	**Mars**	**Venus**
Distance from sun	1 AU	1.5 AU	0.72 AU
$1/\text{Distance}^2$	1.00	0.44	1.9
Solar constant	$1350\ \text{W}/\text{m}^2$	$600\ \text{W}/\text{m}^2$	$2604\ \text{W}/\text{m}^2$
Albedo	0.30	0.17	0.71
$T_{\text{bare rock}}$	254 K (-2°F)	216 K (-70°F)	240 K (-27°F)
T_{surface}	295 K (71°F)	240 K (-28°F)	700 K (800°F)
Δ_T	41 K (74°F)	24 K (42°F)	460 K (828°F)