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Homework 2

All calculations can be found in the notebook
`https://github.com/meredith-durbin/ASTR531/blob/master/HW2/HW2.ipynb`.

8.2 (a) The helium-burning core must be less massive than the hydrogen-burning core for a given stellar mass; for a given temperature/density profile, less of the enclosed mass will be at the temperature required to kick off helium vs. hydrogen burning.

(b) The star will be more chemically stratified at the end of the helium-burning phase, with the helium fusion products at the core and a shell of unfused helium surrounding them.

9.1 Timescales for various stars:

Star	$\tau_{\rm dyn}$	$\tau_{\rm KH}$	$\tau_{\rm nucl}$
MS, 1 $M_\odot$	0.906 h	3.140×10^7 yr	10^{10} yr
MS, 60 $M_\odot$	6.792 h	9.487×10^3 yr	7.554×10^5 yr
RSG, 15 $M_\odot$	5.056 yr	4.793 yr	3.358×10^5 yr
WD, 0.6 $M_\odot$	7.142 s	7.945×10^{10} yr	—

9.2 If nuclear fusion in the sun were to suddenly stop, it would take approximately a thermal timescale to notice; the solar spectrum we observe is largely a product of the thermal timescale is the timescale on which t.

11.2 (a) ???

12.2 (a) All radii are in solar radii.

Mass ($M_\odot$)	$R_{\rm start,H}$	$R_{\rm end,H}$	$R_{\rm end,PMS}$
0.1	10	0.2	0.2
1	100	2	1
10	1000	20	5
100	10^4	200	25

(b) Timescales

Mass ($M_\odot$)	$\tau_{\rm Hayashi}$ (yr)	$\tau_{\rm PMS}$ (yr)
0.1	10^7	1.897×10^{10}
1	10^6	6×10^7
10	10^5	1.897×10^5
100	10^4	6×10^2

13.2 I chose to calculate the ratio of final to initial quantities so that I could ignore mass entirely.

$$\mu_0 = (2 - 1.25Y_0)^{-1} \tag{1}$$

$$\mu_1 = (2 - 1.25Y_1)^{-1} \tag{2}$$

$$L_1/L_0 = \frac{(2 - Y_1)^{-1}\mu_1^4}{(2 - Y_0)^{-1}\mu_1^4} \tag{3}$$

$$R_1/R_0 = \frac{(1 + X_1)^{0.05}\mu_1^{2/3}}{(1 + X_0)^{0.05}\mu_0^{2/3}} \tag{4}$$

$$T_1/T_0 = \frac{(1 + X_1)^{-0.5}\mu_1^{0.83}}{(1 + X_0)^{-0.5}\mu_0^{0.83}} \tag{5}$$

For $X_1 = Y_1 = 0.49$ and $X_0 = 0.7$, $Y_0 = 0.28$, $L_1/L_0 = 2.28$, $R_1/R_0 = 1.12$, and $T_1/T_0 = 1.23$.