

Correction to: STARFORGE: Towards a comprehensive numerical model of star cluster formation and feedback

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13 November 2025

1 THE ERROR

The original STARFORGE simulations methods paper (Grudić et al. 2021) provided a fitting functions for the infrared Planck-mean dust opacity κ_P as a function of dust temperature T_d in Appendix C, derived from the model of Semenov et al. (2003). Although this was the model originally used in the simulations, e.g. Grudić et al. (2022), it is not a good one for this purpose. This is because it neglects the strong dependence of dust opacity on photon energy, or equivalently the radiation temperature T_{rad} defined in the context of our radiative transfer model. Quite often in the ISM $T_{\text{rad}} \neq T_d$, so frequency-integrated opacity must be treated as a function of both temperatures.

Other authors appear to have made a similar error (or omission) while modeling ISM conditions where the distinction between T_{rad} and T_d can be quite important (Dopcke et al. 2011; Smith et al. 2017; Kannan et al. 2020; Deng et al. 2024; Zimmermann et al. 2025). Notably, this nuance was previously pointed out in a footnote in Cunningham et al. (2018), although they did not track the distinct temperatures in their simulation.

We take this opportunity to clarify the use of Planck-mean opacities in the calculation of dust radiative processes out of local thermodynamic equilibrium (LTE), and explicitly compute the dependence of mean opacity on T_{rad} and T_d .

2 PLANCK-MEAN DUST OPACITIES FOR EMISSION AND ABSORPTION

Assuming the heat capacity of dust is very small, the steady-state dust energy equation balances three processes:

$$\underbrace{\int d\nu \kappa_\nu(T_d) \rho c u_\nu}_{\text{Absorption}} - \underbrace{\int d\nu \epsilon_\nu}_{\text{Emission}} + \underbrace{n_H^2 \alpha_{\text{gd}}(T) (T - T_d)}_{\text{Gas-dust collisions}} = 0, \quad (1)$$

where ν is the photon frequency, $\kappa_\nu(T_d)$ is the monochromatic dust absorption opacity, ρ is the mass density, u_ν is the photon energy density per unit frequency, ϵ_ν is the volumetric emissivity, n_H is the number density of H nuclei, T is the gas temperature, and α_{gd} is the gas-dust collision coefficient (Hollenbach & McKee 1989). Here we have made explicit the dependence of $\kappa_\nu(T_d)$ upon dust temperature, due to varying grain composition as volatiles sublime (Semenov et al. 2003).

The assumption we make for the STARFORGE far-IR photon frequency component is

$$u_\nu = u_{\text{IR}} \times \frac{B_\nu(T_{\text{rad}})}{\int d\nu B_\nu(T_{\text{rad}})} \quad (2)$$

i.e. the photon energy distribution is proportional to that of a blackbody with temperature T_{rad} , with frequency-integrated energy density u . Integrating Eq. 1 while neglecting the other radiation bands absorbed by dust:

$$\kappa_P(T_d, T_{\text{rad}}) \rho c u_{\text{IR}} - \epsilon_d + n_H^2 \alpha_{\text{gd}}(T) (T - T_d) = 0, \quad (3)$$

where

$$\kappa_P(T_d, T_{\text{rad}}) = \frac{\int d\nu \kappa_\nu(T_d) B_\nu(T_{\text{rad}})}{\int d\nu B_\nu(T_{\text{rad}})} \quad (4)$$

is the Planck-mean opacity, which has two distinct temperature arguments: the first accounts for variations in dust composition with T_d , and the second dependence upon T_{rad} due to the original frequency-dependence of $\kappa_\nu(T_d)$.

In local thermodynamic equilibrium where $T_d = T_{\text{rad}} = T$ and $u = aT^4$, we require $\epsilon_d = ac\rho T^4 \kappa_P(T, T)$. In general, out of LTE, Kirchoff's law for thermal emission therefore implies that $\epsilon_d = ac\rho \kappa_P(T_d, T_d) T_d^4$. So the frequency-integrated energy balance equation becomes

$$\rho c \left(\kappa_P(T_d, T_{\text{rad}}) u_{\text{IR}} - \kappa_P(T_d, T_d) aT_d^4 \right) + n_H^2 \alpha_{\text{gd}}(T) (T - T_d) = 0. \quad (5)$$

Thus, the same functional form for κ_P is used for both emission and absorption, but the emission term substitutes T_d in the T_{rad} slot while the absorption term uses the distinct temperatures. From it is apparent that $T_{\text{rad}} \sim T_{\text{dust}}$ only under certain conditions, e.g. when the radiative terms are dominant and $u_{\text{IR}} \approx aT_{\text{rad}}^4$. An important counterexample is at high attenuations deep within a pre-stellar molecular cloud, where the ambient optical and UV components are attenuated. Here u_{IR} is dominated by the FIR dust-emission component of the ISM, which is highly diluted compared to a blackbody and hence $T_d \ll T_{\text{rad}}$. This regime of dust energy balance is important for thermal evolution during pre-stellar core collapse (Masunaga et al. 1998; Hennebelle & Grudić 2024).

For completeness, the full equation solved in the current version

of the STARFORGE model for T_d , accounting for all frequency components and radiative processes, is

$$\rho c \left[\sum_i \kappa_i u_i + \kappa_P(T_d, T_{\text{rad}}) u_{\text{IR}} - \kappa_P(T_d, T_d) a T_d^4 \right] + n_H^2 \alpha_{\text{gd}}(T) (T - T_d) = 0, \quad (6)$$

where i runs over the FUV, near-UV, and optical-NIR frequency bands, with corresponding dust opacities κ_i , and $\kappa_{P,g}$ is the Planck-mean opacity of the gas itself to the IR band.

3 MEAN DUST OPACITY AS A FUNCTION OF T_{RAD} AND T_d

We compute the Planck-mean opacity explicitly as a function of T_{rad} and T_d by numerically integrating the monochromatic opacity tables provided by Semenov et al. (2003) for all 5 T_d regions demarcated by sublimation points of the dust components. This routine is implemented in the `radiation.dust` submodule of the `meshoid` Python package¹. However, it is not practical to compute κ_P in this way on-the-fly within the dust temperature solver in GIZMO or another RHD simulation. Instead, we use to a simple log-space linear interpolant of a lookup table as a function of T_{rad} , with a separate table for each of the 5 distinct temperature regions.

For completeness, we also also compute Rosseland-mean opacities (Fig. 2) as a function T_{rad} and T_d . In the radiative diffusion regime where the Rosseland mean is useful, conditions are likely to be closer to LTE and distinction between T_{rad} , T , and T_{dust} is typically less important.

The code used to generate these figures and a set of tabulated mean opacities are available at <https://github.com/mikegrudic/STARFORGE-methods-errata> or <https://doi.org/10.5281/zenodo.17596225>.

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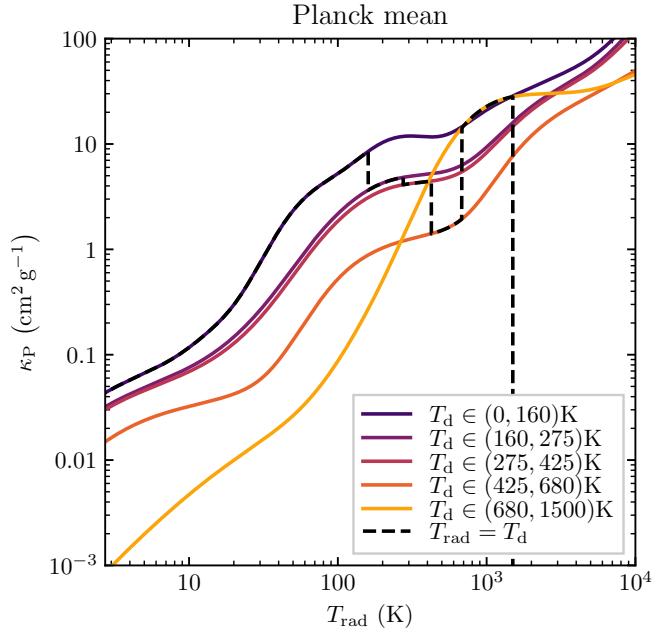


Figure 1. Planck-mean opacity as a function of both dust temperature T_d and radiation temperature T_{rad} , computed from the monochromatic opacity tables of Semenov et al. (2003) for their ‘porous 5-layeredsphere’ model. The dashed line plots the Planck-mean opacity assuming $T_d = T_{\text{rad}}$, which disagrees with $\kappa_P(T_d, T_{\text{rad}})$ in general.

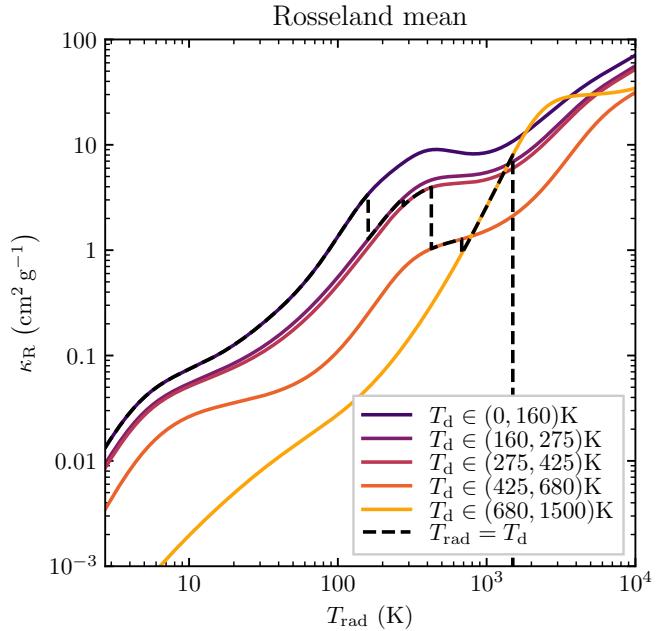


Figure 2. Rosseland mean dust opacity as a function of both dust temperature T_d and radiation temperature T_{rad} , computed from the opacity tables of Semenov et al. (2003) for their ‘porous 5-layeredsphere’ model.

¹ <https://github.com/mikegrudic/meshoid>