

Visible-light Lyot Coronagraph for SCExAO/VAMPIRES

Miles Lucas^{a,*}, Michael Bottom^a, Olivier Guyon^{b,c}, Julien Lozi^b, and Barnaby Norris^d

^aInstitute for Astronomy, University of Hawai'i, 640 N. Aohoku Pl., Hilo, HI 96720, USA

^bNational Observatory of Japan, Subaru Telescope, 650 N. Aohoku Pl., Hilo, HI 96720, USA

^cSteward Observatory, University of Arizona, 933 N. Cherry Ave., Tucson, AZ 85721, USA

^dSydney Institute for Astronomy, School of Physics, Physics Rd., University of Sydney, NSW 2006, Australia

ABSTRACT

We developed a visible-light Lyot coronagraph for SCExAO/VAMPIRES. The coronagraph comprises four hard-edged, partially transmissive focal plane masks with inner working angles of 36 mas, 55 mas, 92 mas and 129 mas, respectively. The Lyot stop is a reflective, undersized design with a relative throughput of 65.7%. Our preliminary raw on-sky contrast from 20 s of median-combined data is 10–2 at 0.1'' to 10–4 at 0.75'' for all mask sizes. The coronagraph was deployed in early 2022 and has been approved for open-use starting in the 2022B semester.

Keywords: Coronagraph, Optical, Visible, High-Contrast, Imaging, Exoplanets

1. INTRODUCTION

In recent decades, the proliferation of high-contrast instruments like SPHERE,¹ GPI,² and SCExAO³ have greatly advanced the field of direct imaging. The workhorse detectors on these instruments typically operate in the near-infrared (near-IR), from Y to L band, for two reasons. First is for the increased brightness of thermal emission of exoplanets in the near-IR, which lowers the relative brightness ratio between the star and the exoplanet (contrast), ultimately improving the detection sensitivity. Second, in the near-IR wavefront control with adaptive optics (AO) is more powerful than the optical, simply due to the longer wavelengths. However, there are significant motivations for extending high-contrast imaging into the optical, particularly for reflected-light detection of exoplanets, polarimetric imaging of protoplanetary disks, and H α emission from accreting stars and exoplanets.

The Visible Aperture-Masking Polarimetric Interferometer/Imager for Resolving Exoplanetary Signatures (VAMPIRES)⁴ is an optical dual-beam imager built into the Subaru Coronagraphic Extreme Adaptive Optics (SCExAO) instrument on the Subaru telescope. VAMPIRES operates between 600 nm to 800 nm with capabilities for polarimetric imaging,⁵ H α imaging,⁶ and focal plane wavefront sensing.⁷ VAMPIRES also utilizes non-redundant sparse aperture masking for sub-diffraction-limited interferometric imaging of stellar surfaces and outflows. For this work, though, we will focus on the traditional and polarimetric imaging modes.

VAMPIRES currently uses two Andor electron-multiplying (EM) CCDs for low readnoise with frame rates from 1 Hz to 100 Hz. While these detectors perform very well in the photon-limited regime, the strong nonlinearities and compressed dynamic range from using the EM registers makes imaging intermediate or bright targets ($m^R < 7$) difficult. In order to address this problem, we have designed and deployed a classic Lyot coronagraph (CLC), which greatly attenuates the stellar point spread function (PSF), allowing use of high EM gain without saturation. Furthermore, coronagraphy is a well-established technique for imaging circumstellar regions by controlling the stellar diffraction pattern. In high-contrast imaging, coronagraphs are an important tool for increasing the detection sensitivity around stars, enabling direct detection and analysis of exoplanets and circumstellar disks. Combined with extreme AO, such as provided by SCExAO, a coronagraph can enable detections of sources many orders of magnitude fainter than their host stars ($\sim 10^{-6}$).⁸

Further author information: (Send correspondence to M.D.L.)

M.D.L.: E-mail: mdlucas@hawaii.edu

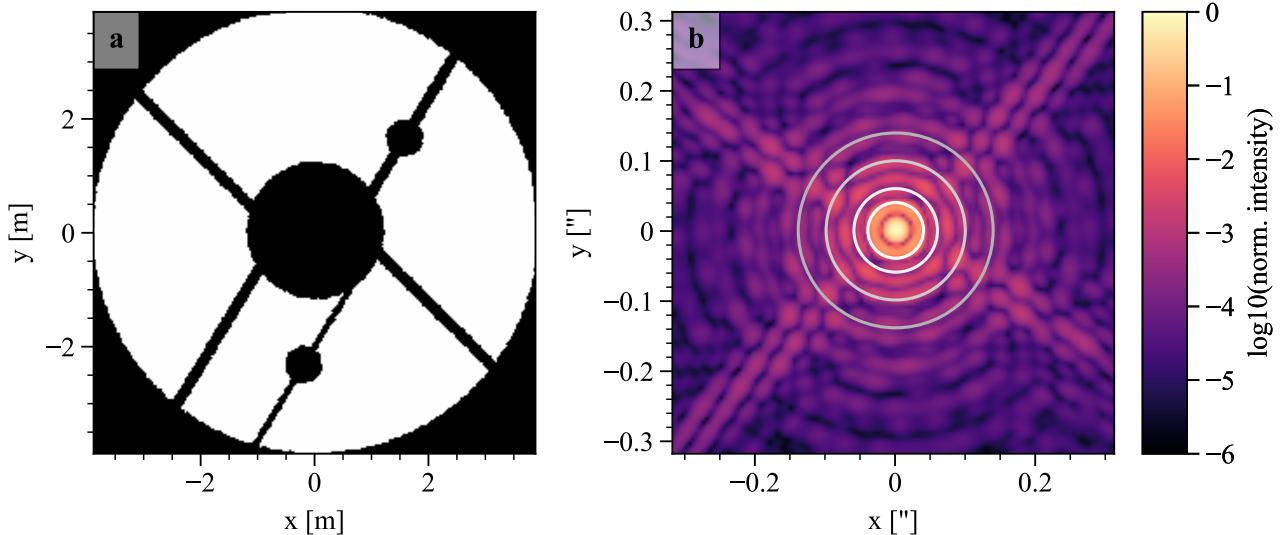


Figure 1. (a) The SCExAO pupil, which is slightly undersized compared to the Subaru aperture and contains additional features. The aperture diameter is 95% of the Subaru clear aperture; effectively 7.79 m. The central obstruction is 31% of this diameter. The two additional circles are masks for inactive deformable mirror actuators, one of which lies on a spider segment, while the other requires an additional spider. (b) The VAMPIRES PSF simulated using HCIPy. The four grayscale circles correspond to the focal plane mask sizes.

This report details our design and deployment of a visible-light coronagraph for VAMPIRES. In section 2 we detail our modeling and simulation of the coronagraph and the construction of the optics. In section 3 we describe our characterization of the coronagraph, including throughput and inner working angle. Finally, in section 4 we discuss the performance of the coronagraph with the calibration light source and on sky.

2. METHODS

2.1 Designing the Coronagraph

We used the open-source python package HCIPy⁹ for simulating our coronagraph design with Fourier optics. We defined our wavefronts on a rasterized image of the SCExAO pupil sampled on a 256x256 pixel grid (Figure 1a). Initially, we do not introduce any wavefront errors. Using a Fraunhofer propagator with an F/28.4 focal ratio, we formed the intermediate focal plane, where the focal plane mask will be inserted. VAMPIRES has a field of view (FOV) of 3''x3'', however, for our exploratory simulations we only calculated the first 16 airy rings (Figure 1b). The focal plane masks are hard-edged circles, and were modeled using a circular apodizer with either 0% or 0.1% transmission. We designed four focal plane masks with radii of 38.6 μm, 57.9 μm, 96.5 μm and 135 μm, respectively. These radii correspond to 2, 3, 5 and 7 λ/D at 750 nm. The focal plane after the mask was propagated to the next pupil plane, where the Lyot stop apodized the wavefront.

The Lyot stop was modeled using a series of circular and rectangular apodizers with 0% transmission. The stop design follows the shape of the SCExAO pupil with an undersized diameter and oversized obstructions. The outer diameter is 90% the clear aperture diameter, the inner obstruction is 45% the clear aperture diameter, the secondary support struts are 200% their original width, and the bad deformable mirror probes are 150% oversized. The ratio of areas of the Lyot stop and the SCExAO pupil is 63.7%, and will theoretically widen the diffraction limit by 5%. This final wavefront was propagated to the detector focal plane, forming the post-coronagraphic point-spread function.

We repeated the above process for a series of wavelengths to approximate a broadband PSF. We summed the images produced with 10 wavefronts from 725 nm to 775 nm to mimic the “750-50” filter (750 nm central wavelength with 50 nm uniform bandpass). We chose this filter because at these wavelengths VAMPIRES has the highest throughput and best polarimetric efficiency. As an additional step, we generated many datasets

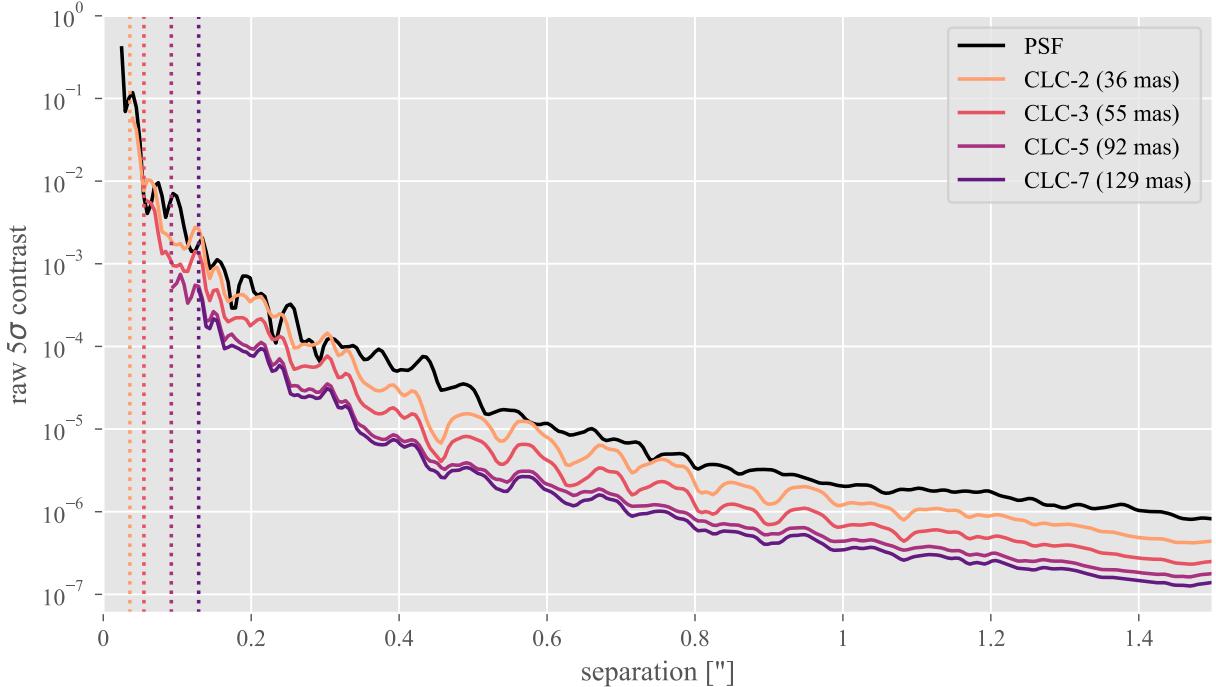


Figure 2. Simulated raw 5σ contrast curves for an unocculted PSF (black line) and for each focal plane mask size (orange to purple lines). These curves are calculated from images simulated with `HCI Py`. The mask inner working angles are marked by vertical dotted lines

with random tip and tilt errors sampled from a bivariate Gaussian with 10 mas root-mean-square (RMS) jitter to test the coronagraph’s resiliency to low-order wavefront errors. Although 10 mas RMS jitter is around the 90th percentile for tip-tilt performance on SCExAO, since VAMPIRES often operates in high-frame-rate lucky imaging mode the effects of tip-tilt can be mitigated.

We evaluated the performance and quality of our simulations from the raw contrast curves of each mask Figure 2. These curves show the 5σ detection thresholds (accounting for small sample statistics¹⁰) without tip-tilt jitter and without any PSF subtraction. For clarity, we describe all our contrast curves as “raw” since we never do PSF subtraction in this work. The curves use a robust scale estimator (biweight middeviance) and are normalized using the flux of the non-coronagraphic PSF. We used the open-source package `ADI.jl`¹¹ for creating these contrast curves.

From the raw contrast curves, we found somewhat poor performance from the CLC-2 and CLC-3 masks compared to the unocculted PSF. We found no significant performance differences between the CLC-5 and CLC-7 masks, perhaps reaching some limit of performance. We note that the typical read noise and non-common path wavefront errors in VAMPIRES were expected to ultimately limit our contrast at large separations, anyways. Due to the convenience, we decided to construct all four mask sizes with two variations each: one partial transmissive (0.1%) and the other fully opaque.

2.2 Construction and Installation

We chose a metallic vapor deposition process (<https://opto-line.com/>) for both the focal plane masks and the Lyot stop for its consistency and precision for custom patterns. The metals are deposited onto a glass substrate with micrometer-precision, giving us great flexibility and precision in our designs. One downside of this approach is that the glass focal plane mask will shift the focus since it is located in a converging beam. In VAMPIRES, there is a custom optics mount in the focal plane with 8 mm x 8 mm slots, so we arranged for four masks to be diced from a single 30 mm anti-reflection (AR) coated optical flat ([Edmund Optics #84-466](#)), making sure each segment was within the clear aperture. The focal plane mask dots were made using chromium

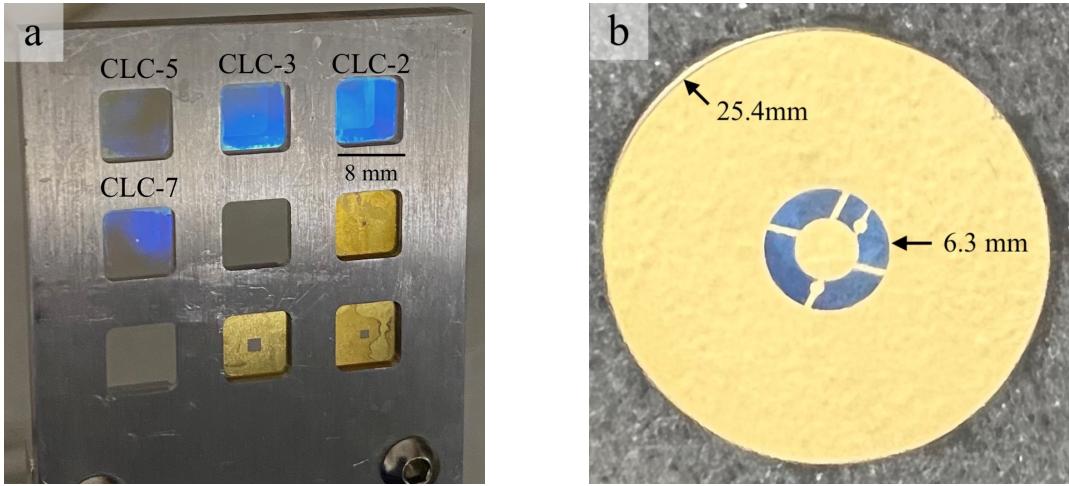


Figure 3. (a) The four focal plane masks mounted in the custom-machined focal plane mount for VAMPIRES. The size of each slot is 8 mm x 8 mm. (b) The undersized Lyot stop with its reflective gold coating. The diameter of the glass substrate and stop outer diameter are marked.

with a thicknesses of 110 nm and 300 nm. The thickness was determined from the optical penetration depth for a desired transmission

$$\hat{s}(\lambda) = -\frac{\lambda}{4\pi\tilde{k}(\lambda)} \ln \frac{I(\lambda)}{I_0(\lambda)} \quad (1)$$

where \hat{s} is the thickness, λ is the wavelength (tested across the 600-800 nm range for VAMPIRES), \tilde{k} is the extinction coefficient, and I/I_0 is the relative intensity for transmitted light. The thicknesses of 110 nm and 300 nm correspond to transmissions of 1e-1 and 1e-8, respectively, matching our design goal of 0.1% and 0%.

The Lyot stop pattern was deposited onto a 25 mm AR-coated optical flat ([Edmund Optics #84-465](#)) with no further processing. We designed the Lyot stop to be reflective for the purposes of alignment and low-order wavefront sensing from the rejected light. We chose gold for the surface material for its high reflectivity (>90%) across the 600 nm to 800 nm wavelength range. Following advice from our vendor, this gold layer was deposited on top of a layer of chromium, which adheres better to AR coatings than gold. We used [Equation 1](#) with a linear combination of 50% chromium and 50% gold to determine a combined thickness of 285 nm for 1e-8 transmission. The final thickness of the chromium and gold deposit was specified as 300 nm.

The focal plane masks and Lyot stop were installed on SCExAO in January, 2022. We decided to only install the four transmissive (0.1%) focal plane masks given the limited number of slots in the focal plane mount ([Figure 3](#)). The Lyot stop pattern was accidentally mirrored with respect to the beam and was not correctly scaled to the beam width. A new stop was designed and procured to address these issues and was deployed on SCExAO in February, 2022.

3. OPERATION AND CHARACTERIZATION

Both the focal plane masks and Lyot stop are installed in mounts with precise motor controllers for alignment. We found that the Lyot stop, being in the collimated beam, stays consistently aligned in-between instrument cranes. The focal plane masks take around a minute to manually align using a text user interface (TUI) on a high signal to noise ratio target. In addition, because the focal plane masks puts glass in a converging beam, the focus of the instrument shifts and an automatic focus measurement is performed with an offset focal plane mask before use.

3.1 Artificial Calibration Speckles (“Astrogrid”)

A common problem with coronagraphic imaging is accurately determining the flux and position of the host star while behind the focal plane mask. By adding sinusoidal patterns to the deformable mirror (DM) copies of

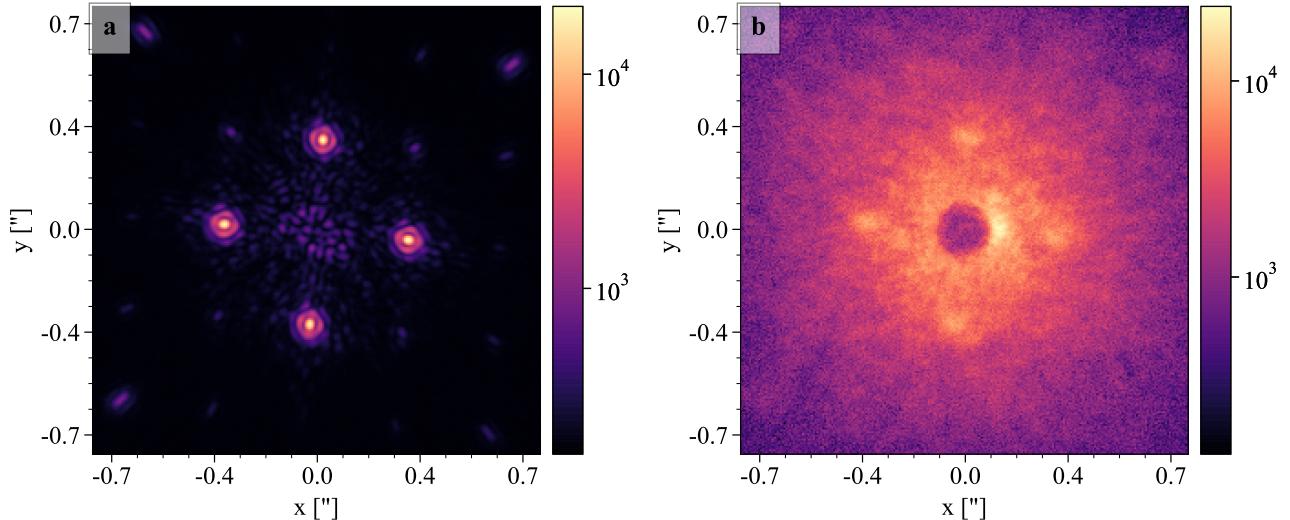


Figure 4. Coronagraphic images using the CLC-5 mask (91 mas IWA) with 50 nm “astrogrid” calibration speckles. (a) A single frame of testbed data using the internal calibration source through the 750-50 filter with 110 EM gain, and 20 ms exposure length. (b) On sky data of HIP 56083 from May 25, 2022 in poor seeing ($\sim 1.6''$), using the 750-50 filter, 300 EM gain, and 200 ms exposure length.

the stellar PSF will appear off-axis depending on the frequency of the sinusoidal pattern.¹² These patterns are modulated in two ways: first the patterns are modulated spatially by half a wavelength to average effects of atmospheric seeing on the speckles during a typical integration. Second, the pattern is spatially modulated to create alternating sets of two PSFs, for a total of four calibration speckles equidistant from the star. The brightness of these “satellite spots” is proportional to the amplitude of the DM perturbation, which can be calibrated to give accurate relative photometry between the spots and the star. This method has been successfully demonstrated on SCExAO/CHARIS with astrometric precision ~ 1.7 mas and photometric precision $\sim 0.3\%$.¹³ Examples of the “astrogrid” on VAMPIRES are shown in Figure 4.

3.2 Coronagraphic throughput

We measured both the relative Lyot stop throughput as well as the off-axis focal plane throughput for our coronagraph. The Lyot stop throughput is proportional to the unobstructed area of the mask compared to the pupil. From the drawings, the ratio of areas is 63.7%. We measured the throughput using the pupil-imaging camera with the calibration light source, comparing the total intensity measured with a mirror in the pupil wheel (100% of the incident light) versus the Lyot stop (blocked light only). We calculated a relative throughput of 65.7%.

The off-axis coronagraphic throughput was measured by gradually moving the focal plane mask off-axis from the calibration source. The fluxes from all of the images were normalized so the maximum off-axis flux was 1 and the on-axis flux was 0 (Figure 5). Then, these curves were fit with a logistic function, allowing precise determination of the inner working angle (IWA), where the throughput is 50%. The IWAs for each mask were measured to be 36 mas, 55 mas, 92 mas and 129 mas, respectively, which is in good agreement with the radii of the masks.

4. RESULTS

4.1 Testbed results

Our first tests use the calibration source on SCExAO for determining the contrast of the coronagraph for each focal plane mask. For these tests we used the 750-50 filter since it has the highest throughput. In addition, for all tests we applied an astrogrid with 50 nm amplitude and a spatial frequency of $15.5 \lambda/D$ (308 mas).

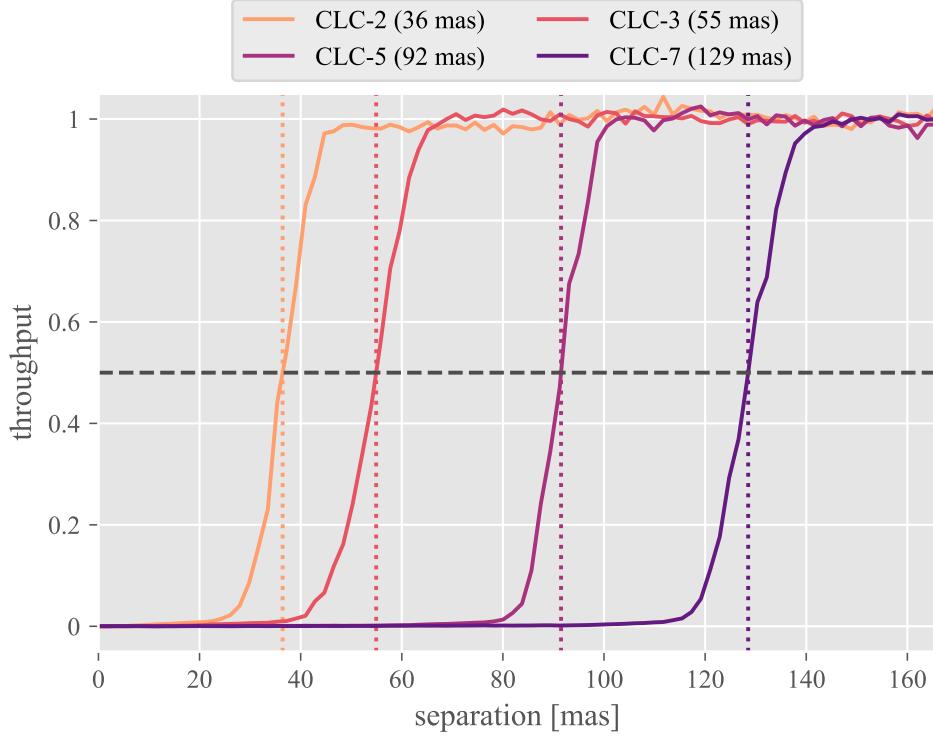


Figure 5. Off-axis coronagraphic throughput for each focal plane mask. The flux throughput is normalized between 0 and 1. 50% throughput is marked with a black dashed line, and the inner working angles, fit using a logistic function (light gray lines), are denoted by the orange to purple dotted lines.

First, we took non-coronagraphic data allowing us to characterize both the photometry and astrometry of the astrogrid speckles as well as giving a non-coronagraphic PSF for measuring the relative attenuation of the coronagraph. For data reduction, we median-combined a series of 100 images and recentered the frame using the satellite spot centers of mass. For the astrogrid calibration, we used a hierarchical model of five two-dimensional Moffat functions

$$f(x, y | x_0, y_0, \Gamma, \alpha, I_0) = I_0 \left[1 + \frac{(x - x_0)^2 + (y - y_0)^2}{\Gamma^2} \right]^{-\alpha} \quad (2)$$

where (x_0, y_0) is the center of the PSF, I_0 is the peak amplitude, Γ is the half-width at half-maximum (HWHM), and α is the exponential scaling factor, typically around 1 to 3. In our hierarchical model each PSF shares the Γ and α parameters, and each of the four satellite spots share an amplitude. Around each PSF we measured the photometry using circular apertures with background subtraction from annular apertures. We measured the relative photometry of the 50 nm amplitude astrogrid as $10^{-1.52}$.

Then we inserted the focal plane mask and Lyot stop and refocused. Again, we median-combined a series of 100 images and recentered using the satellite spots. To measure the photometry we fit a hierarchical model of four two-dimensional Moffat functions (Equation 2). This follows the same procedure as above with shared shape parameters and intensities, but this time without the on-axis PSF. Then, we measured the photometry of each satellite spot using circular apertures with an annulus for background subtraction. The stellar flux is estimated by the median flux of the four satellite spots scaled by the relative photometry fit above ($10^{-1.52}$).

From here we create raw contrast curves the same way as in subsection 2.1, calculating the 5σ detection limits (accounting for small-sample statistics) without PSF subtraction. Because we are not performing any kind of differential imaging, the satellite spots heavily bias the contrast, so we masked out the satellite spots during the noise calculations.

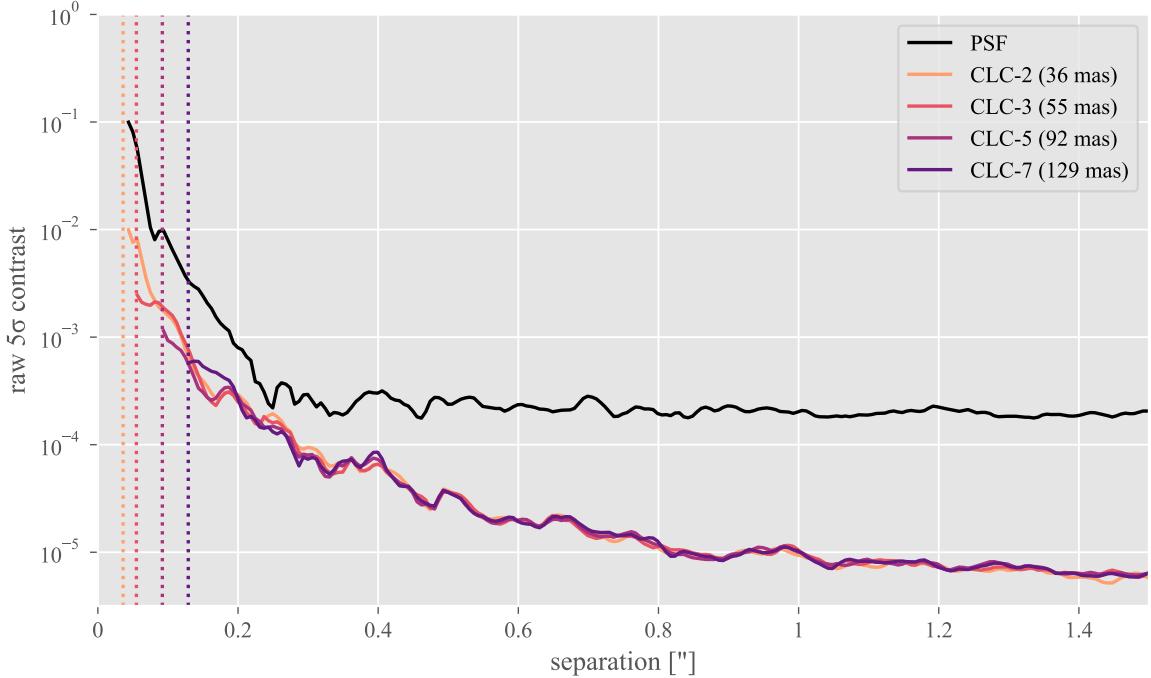


Figure 6. Raw 5σ contrast curves measured from 2s of testbed data. These curves account for small-sample statistics and have no PSF subtraction. The black curve is the non-coronagraphic contrast, while the orange through purple curves correspond to each focal plane mask. The IWA for each curve is marked with a dotted vertical line. The astrogrid speckles are masked out during the calculation, but their position is marked with a gray vertical line.

From the raw contrast curves (Figure 6), we note similar performance between all the mask sizes. At $0.1''$, the coronagraph achieves a raw contrast of $\sim 10^{-3}$, and a read-noise limited contrast beyond $0.8''$ of $\sim 10^{-5}$. There is a large bump around the location of the satellite spots ($\sim 0.3''$) and, to a lesser extent, around the secondary spots ($\sqrt{2} \cdot r \sim 0.42''$). Even with the masking, these spots are biased, we believe due to speckle pinning.¹⁴ We note that throughout the entire FOV the coronagraphs consistently attenuate the light by 1 to 1.5 orders of magnitude, addressing the dynamic range problems currently faced by VAMPIRES.

4.2 On-sky results

We have not had any opportunities to use the coronagraph in good conditions for testing; most of our on-sky tests had $>1.5''$ seeing. Our best data is from May 12, 2022 of target HIP 56083. We used a 50 nm amplitude astrogrid with a separation of $15.5 \lambda/D$, the same as our testbed data. An example coronagraphic image is shown in Figure 4b. We used 300 EM gain with 200 ms exposure times, median combining 100 frames for a total of 20s of integration time. This caused significant blurring from seeing, but was necessary to achieve a high signal to noise ratio (S/N). In the future, we want to investigate further the trade-off between S/N and sharpness from shorter exposure times on coronagraphic contrast.

We constructed 5σ contrast curves for each focal plane mask size following the same procedure as the testbed data (Figure 7). Note for these observations we used a tighter field of view (256x256) which limits the max separation to $0.75''$. We used the same scaling for the satellite spot photometry as fit from the testbed data ($10^{-1.52}$) and also masked out the satellite spots. At $0.1''$, all masks have a raw contrast around $10^{-2.5}$ and drop off to 10^{-4} at $0.75''$. We expect that under better observing conditions we could achieve one to two orders of magnitude increased contrast performance with high S/N data aligned and stacked using lucky imaging for improved sharpness.

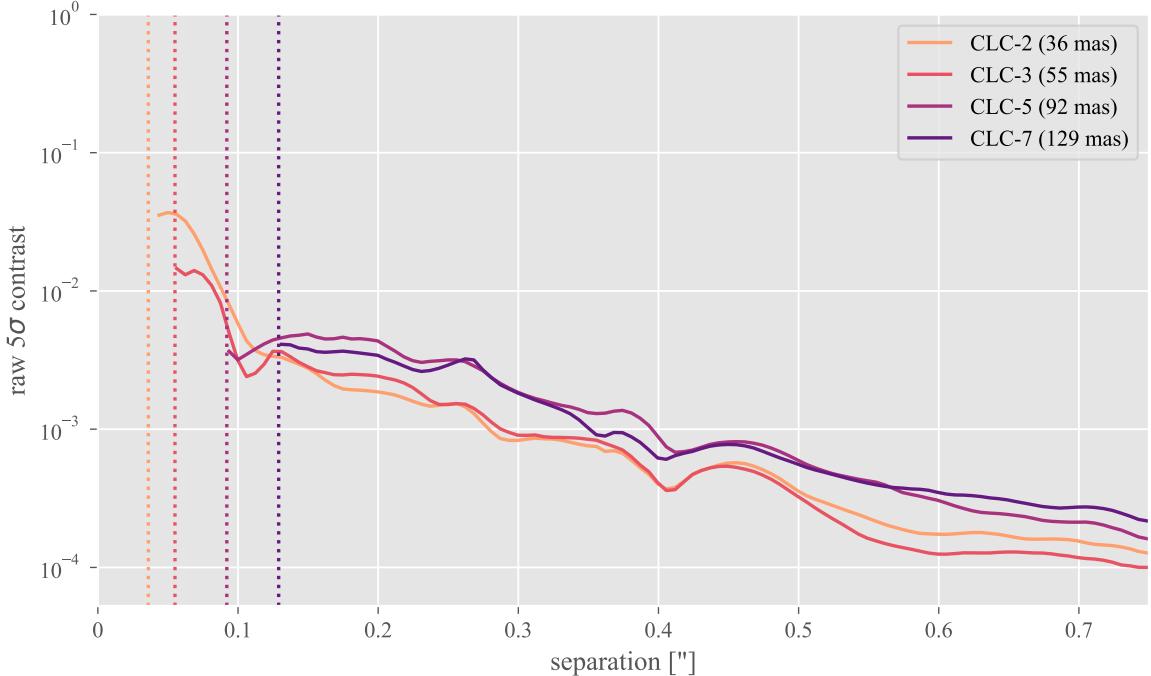


Figure 7. Raw 5σ contrast curves measured from 20 s of on-sky data of HIP 56083. These data were obtained on May 12, 2022 through the 750-50 filter with 300 EM gain and 200 ms exposures. The orange through purple curves correspond to each focal plane mask. The IWA for each curve is marked with a dotted vertical line. The astrogrid speckles are masked out during the calculation, but their position is marked with a gray vertical line.

5. CONCLUSIONS

In this paper we have described the design, construction, and deployment of a visible-light Lyot coronagraph on SCExAO/VAMPIRES. This coronagraph attenuates 1 to 1.5 orders of magnitude of light, greatly improving the dynamic range of the EM-CCDs used on VAMPIRES with max EM gain. Additionally, the diffraction control will improve high-contrast performance which has important applications in the optical for exoplanet science, both companion and protoplanetary disk imaging, as well as H α imaging. The coronagraph was approved for open-use observation starting in the 2022B observing semester.

Four partially transmissive focal plane masks were installed with IWAs of 36 mas, 55 mas, 92 mas and 129 mas, respectively. The Lyot stop has a measured throughput of 65.7%. During operation, the coronagraph uses “astrogrid” artificial speckles for photometric and astrometric calibration. We tested the coronagraph using the SCExAO internal source for characterizing the IWA, Lyot stop throughput, satellite spot relative flux, and measured contrast curves for each mask size using 2 s of median-combined data. The contrast was similar for each mask size, spanning 10^{-3} to 10^{-5} from $0.1''$ to $>0.8''$. We also measured the contrast for each mask size using 20 s of median-combined data of HIP 56083 in poor seeing conditions. This data is highly limited by the seeing halo, but achieves raw 5σ contrast of 10^{-2} to 10^{-4} from $0.1''$ to $0.75''$.

In the future, as wavefront control improves for SCExAO, we want to explore phase-based focal plane masks such as the four-quadrant phase mask (4QPM).¹⁵ Phase-based masks offer improved contrast over amplitude-based designs like the classic Lyot style, but are highly sensitive to pointing errors and thus require fine tip-tilt control for effective use.¹⁶ The fine tip-tilt control is expected to improve on SCExAO during the upgrades to the facility AO system, AO 188. We are also interested in exploring apodized Lyot stops to achieve similar or improved performance in the presence of aberrations without sacrificing as much throughput.

APPENDIX A. CODE AND DATA AVAILABILITY

The code used for simulating the coronagraph ([subsection 2.1](#)), reducing and analyzing testbed ([subsection 4.1](#)) and on-sky data ([subsection 4.2](#)), and for producing the figures in this paper are all available under an open-source license in a GitHub repository ([mileslucas/vampires-coronagraph](#)). This code makes significant use of the Julia programming language¹⁷ along with the open-source packages `HCIPy`,⁹ `ADI.jl`,¹¹ `numpy`,¹⁸ and `astropy`.^{19,20} The CAD drawings for the focal plane masks and Lyot stops are also available in the online repository. Further requests or questions about code or data are welcomed.

ACKNOWLEDGMENTS

We wish to recognize and acknowledge the significant cultural role and reverence that the summit of Maunakea has always had within the indigenous Hawaiian community. We are most fortunate and thank the community for the privilege to conduct observations from this mountain. We also would like to thank Paul Sumner for his useful advice during the construction of the Lyot stop and focal plane masks. Lastly, we thank Jonathan Williams, Thayne Currie, and Timothy Brandt for generously offering observation time to use VAMPIRES for additional characterization of the coronagraph on-sky. This work is based on data collected at Subaru Telescope, which is operated by the National Astronomical Observatory of Japan. The development of SCExAO was supported by the Japan Society for the Promotion of Science (Grant-in-Aid for Research #23340051, #26220704, #23103002, #19H00703, and #19H00695), the Astrobiology Center of the National Institutes of Natural Sciences, Japan, the Mt Cuba Foundation and the director’s contingency fund at Subaru Telescope. This research was funded by the Heising-Simons Foundation through grant #2020-1823.

REFERENCES

- [1] Petit, C., Sauvage, J.-F., Fusco, T., Sevin, A., Suarez, M., Costille, A., Vigan, A., Soenke, C., Perret, D., Rochat, S., Barrufolo, A., Salasnich, B., Beuzit, J.-L., Dohlen, K., Mouillet, D., Puget, P., Wildi, F., Kasper, M., Conan, J.-M., Kulcsár, C., and Raynaud, H.-F., “SPHERE eXtreme AO control scheme: Final performance assessment and on sky validation of the first auto-tuned LQG based operational system,” in [*Adaptive Optics Systems IV*], **9148**, 91480O, International Society for Optics and Photonics (Aug. 2014).
- [2] Macintosh, B., Graham, J. R., Ingraham, P., Konopacky, Q., Marois, C., Perrin, M., Poyneer, L., Bauman, B., Barman, T., Burrows, A. S., Cardwell, A., Chilcote, J., Rosa, R. J. D., Dillon, D., Doyon, R., Dunn, J., Erikson, D., Fitzgerald, M. P., Gavel, D., Goodsell, S., Hartung, M., Hibon, P., Kalas, P., Larkin, J., Maire, J., Marchis, F., Marley, M. S., McBride, J., Millar-Blanchaer, M., Morzinski, K., Norton, A., Oppenheimer, B. R., Palmer, D., Patience, J., Pueyo, L., Rantakyro, F., Sadakuni, N., Saddlemyer, L., Savransky, D., Serio, A., Soummer, R., Sivaramakrishnan, A., Song, I., Thomas, S., Wallace, J. K., Wiktorowicz, S., and Wolff, S., “First light of the Gemini Planet Imager,” *PNAS* **111**, 12661–12666 (Sept. 2014).
- [3] Jovanovic, N., Martinache, F., Guyon, O., Clergeon, C., Singh, G., Kudo, T., Garrel, V., Newman, K., Doughty, D., Lozi, J., Males, J., Minowa, Y., Hayano, Y., Takato, N., Morino, J., Kuhn, J., Serabyn, E., Norris, B., Tuthill, P., Schworer, G., Stewart, P., Close, L., Huby, E., Perrin, G., Lacour, S., Gauchet, L., Viard, S., Murakami, N., Oshiyama, F., Baba, N., Matsuo, T., Nishikawa, J., Tamura, M., Lai, O., Marchis, F., Duchene, G., Kotani, T., and Woillez, J., “The Subaru Coronagraphic Extreme Adaptive Optics System: Enabling High-Contrast Imaging on Solar-System Scales,” *Publications of the Astronomical Society of the Pacific* **127**, 890 (Sept. 2015).
- [4] Norris, B., Schworer, G., Tuthill, P., Jovanovic, N., Guyon, O., Stewart, P., and Martinache, F., “The VAMPIRES instrument: Imaging the innermost regions of protoplanetary discs with polarimetric interferometry,” *Monthly Notices of the Royal Astronomical Society* **447**, 2894–2906 (Mar. 2015).
- [5] Norris, B. R. M., Tuthill, P., Jovanovic, N., Lozi, J., Guyon, O., Cvetojevic, N., and Martinache, F., “Diffraction-limited polarimetric imaging of protoplanetary disks and mass-loss shells with VAMPIRES,” in [*Advances in Optical Astronomical Instrumentation 2019*], **11203**, 112030S, International Society for Optics and Photonics (Jan. 2020).
- [6] Uyama, T., Norris, B., Jovanovic, N., Lozi, J., Tuthill, P., Guyon, O., Kudo, T., Hashimoto, J., Tamura, M., and Martinache, F., “High-contrast H α imaging with Subaru/SCExAO + VAMPIRES,” *Journal of Astronomical Telescopes, Instruments, and Systems* **6**, 045004 (Oct. 2020).

- [7] Vieuard, S., Bos, S. P., Cassaing, F., Currie, T., Deo, V., Guyon, O., Jovanovic, N., Keller, C. U., Lamb, M., Lopez, C., Lozi, J., Martinache, F., Miller, K., Montmerle-Bonnefois, A., Mugnier, L. M., N'Diaye, M., Norris, B., Sahoo, A., Sauvage, J.-F., Skaf, N., Snik, F., Wilby, M. J., and Wong, A., “Focal plane wavefront sensing on SUBARU/SCEExAO,” in [*Adaptive Optics Systems VII*], **11448**, 1261–1277, SPIE (Dec. 2020).
- [8] Guyon, O., “Extreme Adaptive Optics,” *Annu. Rev. Astron. Astrophys.* **56**, 315–355 (Sept. 2018).
- [9] Por, E. H., Haffert, S. Y., Radhakrishnan, V. M., Doelman, D. S., van Kooten, M., and Bos, S. P., “High Contrast Imaging for Python (HCIPy): An open-source adaptive optics and coronagraph simulator,” **10703**, 1070342 (July 2018).
- [10] Mawet, D., Milli, J., Wahhaj, Z., Pelat, D., Absil, O., Delacroix, C., Boccaletti, A., Kasper, M., Kenworthy, M., Marois, C., Mennesson, B., and Pueyo, L., “Fundamental limitations of high contrast imaging set by small sample statistics,” *ApJ* **792**, 97 (Aug. 2014).
- [11] Lucas, M. and Bottom, M., “ADI.jl: A Julia Package for High-Contrast Imaging,” *Journal of Open Source Software* **5**, 2843 (Dec. 2020).
- [12] Sahoo, A., Guyon, O., Lozi, J., Chilcote, J., Jovanovic, N., Brandt, T., Groff, T., and Martinache, F., “Precision Photometric and Astrometric Calibration Using Alternating Satellite Speckles,” *The Astronomical Journal* **159**, 250 (June 2020).
- [13] Currie, T., Guyon, O., Lozi, J., Sahoo, A., Vieuard, S., Deo, V., Chilcote, J., Groff, T., Brandt, T. D., Lawson, K., Skaf, N., Martinache, F., and Kasdin, N. J., “On-sky performance and recent results from the Subaru coronagraphic extreme adaptive optics system,” **11448**, 114487H (Dec. 2020).
- [14] Bloemhof, E. E., “Remnant Speckles in a Highly Corrected Coronagraph,” *The Astrophysical Journal* **610**, L69–L72 (July 2004).
- [15] Rouan, D., Baudrand, J., Boccaletti, A., Baudoz, P., Mawet, D., and Riaud, P., “The Four Quadrant Phase Mask Coronagraph and its avatars,” *Comptes Rendus Physique* **8**, 298–311 (Apr. 2007).
- [16] Huby, E., Bottom, M., Femenia, B., Ngo, H., Mawet, D., Serabyn, E., and Absil, O., “On-sky performance of the QACITS pointing control technique with the Keck/NIRC2 vortex coronagraph,” *Astronomy and Astrophysics* **600**, A46 (Apr. 2017).
- [17] Bezanson, J., Edelman, A., Karpinski, S., and Shah, V. B., “Julia: A Fresh Approach to Numerical Computing,” *SIAM Rev.* **59**, 65–98 (Jan. 2017).
- [18] Harris, C. R., Millman, K. J., van der Walt, S. J., Gommers, R., Virtanen, P., Cournapeau, D., Wieser, E., Taylor, J., Berg, S., Smith, N. J., Kern, R., Picus, M., Hoyer, S., van Kerkwijk, M. H., Brett, M., Haldane, A., del Río, J. F., Wiebe, M., Peterson, P., Gérard-Marchant, P., Sheppard, K., Reddy, T., Weckesser, W., Abbasi, H., Gohlke, C., and Oliphant, T. E., “Array programming with NumPy,” *Nature* **585**, 357–362 (Sept. 2020).
- [19] Collaboration, A., Robitaille, T. P., Tollerud, E. J., Greenfield, P., Droettboom, M., Bray, E., Aldcroft, T., Davis, M., Ginsburg, A., Price-Whelan, A. M., Kerzendorf, W. E., Conley, A., Crighton, N., Barbary, K., Muna, D., Ferguson, H., Grollier, F., Parikh, M. M., Nair, P. H., Unther, H. M., Deil, C., Woillez, J., Conseil, S., Kramer, R., Turner, J. E. H., Singer, L., Fox, R., Weaver, B. A., Zabalza, V., Edwards, Z. I., Azalee Bostroem, K., Burke, D. J., Casey, A. R., Crawford, S. M., Dencheva, N., Ely, J., Jenness, T., Labrie, K., Lim, P. L., Pierfederici, F., Pontzen, A., Ptak, A., Refsdal, B., Servillat, M., and Streicher, O., “Astropy: A community Python package for astronomy,” *Astronomy and Astrophysics* **558**, A33 (Oct. 2013).
- [20] Astropy Collaboration, Price-Whelan, A. M., Sipőcz, B. M., Günther, H. M., Lim, P. L., Crawford, S. M., Conseil, S., Shupe, D. L., Craig, M. W., Dencheva, N., Ginsburg, A., VanderPlas, J. T., Bradley, L. D., Pérez-Suárez, D., de Val-Borro, M., Aldcroft, T. L., Cruz, K. L., Robitaille, T. P., Tollerud, E. J., Ardelean, C., Babej, T., Bach, Y. P., Bachetti, M., Bakanov, A. V., Bamford, S. P., Barentsen, G., Barmby, P., Baumbach, A., Berry, K. L., Biscani, F., Boquien, M., Bostroem, K. A., Bouma, L. G., Brammer, G. B., Bray, E. M., Breytenbach, H., Buddelmeijer, H., Burke, D. J., Calderone, G., Cano Rodríguez, J. L., Cara, M., Cardoso, J. V. M., Cheedella, S., Copin, Y., Corrales, L., Crichton, D., D’Avella, D., Deil, C., Depagne, É., Dietrich, J. P., Donath, A., Droettboom, M., Earl, N., Erben, T., Fabbro, S., Ferreira, L. A., Finethy, T., Fox, R. T., Garrison, L. H., Gibbons, S. L. J., Goldstein, D. A., Gommers, R., Greco, J. P., Greenfield, P., Groener, A. M., Grollier, F., Hagen, A., Hirst, P., Homeier, D., Horton, A. J., Hosseinzadeh, G., Hu, L.,

Hunkeler, J. S., Ivezić, Ž., Jain, A., Jenness, T., Kanarek, G., Kendrew, S., Kern, N. S., Kerzendorf, W. E., Khvalko, A., King, J., Kirkby, D., Kulkarni, A. M., Kumar, A., Lee, A., Lenz, D., Littlefair, S. P., Ma, Z., Macleod, D. M., Mastropietro, M., McCully, C., Montagnac, S., Morris, B. M., Mueller, M., Mumford, S. J., Muna, D., Murphy, N. A., Nelson, S., Nguyen, G. H., Ninan, J. P., Nöthe, M., Ogaz, S., Oh, S., Parejko, J. K., Parley, N., Pascual, S., Patil, R., Patil, A. A., Plunkett, A. L., Prochaska, J. X., Rastogi, T., Reddy Janga, V., Sabater, J., Sakurikar, P., Seifert, M., Sherbert, L. E., Sherwood-Taylor, H., Shih, A. Y., Sick, J., Silbiger, M. T., Singanamalla, S., Singer, L. P., Sladen, P. H., Sooley, K. A., Sornarajah, S., Streicher, O., Teuben, P., Thomas, S. W., Tremblay, G. R., Turner, J. E. H., Terrón, V., van Kerkwijk, M. H., de la Vega, A., Watkins, L. L., Weaver, B. A., Whitmore, J. B., Woillez, J., Zabalza, V., and Astropy Contributors, “The Astropy Project: Building an Open-science Project and Status of the v2.0 Core Package,” *The Astronomical Journal* **156**, 123 (Sept. 2018).