



ACYCLE

a time-series analysis software for paleoclimate projects

version 0.2.5

User's Guide

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1. Copyrights

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Any questions regarding the license or the operation of this software may be directed to:

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2. References

If you publish any result generated by this program, please cite the following reference:

- Li, Mingsong, Hinnov, Linda, Kump, Lee. *Acycle: a time-series analysis software for education and paleoclimate projects*, under review, *Computers and Geosciences*

If you publish results using correlation coefficient (**COCO** or **eCOCO**) method, please **also** cite the following references:

- Li, Mingsong, Kump, Lee, Hinnov, Linda, Mann, Michael, 2018. Tracking variable sedimentation rates and astronomical forcing in Phanerozoic paleoclimate proxy series with evolutionary correlation coefficients and hypothesis testing. *Earth and Planetary Science Letters* 501, 165-179.

If you publish results using sedimentary noise model (**DYNOT** or ρ_1 methods), please **also** cite this paper:

- Li, Mingsong, Hinnov, Linda, Huang, Chunju, Ogg, James, 2018. Sedimentary noise and sea levels linked to land–ocean water exchange and obliquity forcing. *Nature communications* 9, 1004. Doi: 10.1038/s41467-018-03454-y

If you publish results using power decomposition analysis (**pda.m**), please **also** cite this paper:

- Li, Mingsong, Huang, Chunju, Hinnov, Linda, Ogg, James, Chen, Zhong-Qiang, Zhang, Yang, 2016. Obliquity-forced climate during the Early Triassic hothouse in China. *Geology* 44, 623-626. doi: 10.1130/G37970.1

If you publish results using **evolutionary FFT** method, please **also** cite this book:

- Kodama, K.P., Hinnov, L., 2015. *Rock Magnetic Cyclostratigraphy*. Wiley-Blackwell.

If you publish results using **Wavelet analysis** method, please **also** cite this paper:

- Torrence, C., Compo, G.P., 1998. A practical guide to wavelet analysis. Bulletin of the American Meteorological society 79, 61-78..

If you publish results using **Bayesian Changepoint** method, please **also** cite this paper:

- Ruggieri, E., 2013. A Bayesian approach to detecting change points in climatic records. International Journal of Climatology 33, 520-528.

If you publish results using **Gaussian processes**, please **also** cite this paper:

- Rasmussen, C.E., Nickisch, H., 2010. Gaussian processes for machine learning (GPML) toolbox. Journal of machine learning research 11, 3011-3015.

3. Software Specifications

3.1 System Requirements

This software is developed in MatLab version 2015b. It was tested in Mac OS El Capitan system.

[1. MatLab version for both Mac and Windows]: Recommended. MatLab is essential for the Acycle package. The package works with both Mac and Windows OS.

[2. Windows version]: This software is a stand-alone program. It was tested in Windows 10 OS. If the computer runs with no MatLab, MatLab runtime is essential for the Acycle stand-alone software. MatLab runtime can be downloaded at: www.matlab.com.

[3. Mac version]: This software is a stand-alone program. It was tested in Mac OS El Capitan system. If the Mac runs with no MatLab, MatLab runtime is essential for the Acycle stand-alone software. MatLab runtime can be downloaded at: www.matlab.com.

3.2 Downloading the ACycle software

The ACycle software is available for download from

GitHub (<https://github.com/mingsongli/acycle/>), **Dropbox** (<https://www.dropbox.com/sh/t53vjs539gmxnm/AAC0BqTR0U5xghKwuVc1Iwbma?dl=0>), or **Baidu Cloud** (https://pan.baidu.com/s/14-xRzV_-BBrE6XfyR_71Nw).

3.3 Installation

[1. MatLab version]: Unzip the Acycle software to your root directory. No installation is needed.

[Warning: if the acycle folder is not in your root directory, users need to make sure the directory contains no SPACE or no language other than ENGLISH.]

[2. Windows version]: *coming ...*

[3. Mac version]: *coming ...*

3.4 Startup

[1. MatLab version]:

- (1) Startup MatLab
- (2) Change the MatLab working directory to the acycle directory
- (3) Type

ac

in MatLab's command window, then press the **Enter** key.

3.5 Data requirement

The input file of data series can be in a variety of formats, including comma-, table- or space-delimited text (*.txt), and comma-separated values files (*.csv) from an Excel-type spreadsheet.

The most data file must contain two columns of series. The first column must be in depth or time, and the second column should be value in the corresponding depth or time.

The data can be saved in any directory and is recommended to save in *Acycle* “data” folder. All data files, plots, and folders are displayed in the GUI list box (Fig. 1).

4. Acycle GUI

4.1 Functions and GUI

Acycle contains the following functions.

File

(New Folder; New Text File; Save *.AC.fig; Open *.fig File; Open Working Directory; Extract Data)

Edit

(Rename; Cut; Copy; Paste; Delete)

Plot

(Plot; Plot PLUS; Plot Standardized; Plot Swap Axis; Stairs, Sampling Rate; Data Distribution; ECOCO Plot)

Basic Series

(Insolation; Astronomical Solution; LR04 Stack; Sine Wave; White Noise; Red Noise)

Math

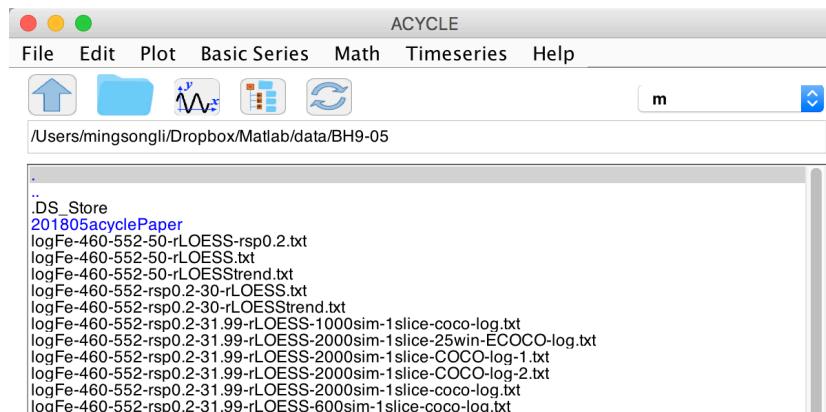
(Sort/Unique/Delete-empty; Select Parts; Merge Series; Add Gaps; Remove Parts; Remove Peaks; Clipping, Interpolation; Smoothing[Moving Average, Bootstrap, Gauss Process]; Changepoint; Sampling Rate Sensitivity; Standardize; Principal Component; Log-transform; First Difference; Derivative; Simple Function; Utilities[Find Max/Min]; Image[Show Image, RGB to Grayscale; Image Profile])

Time series

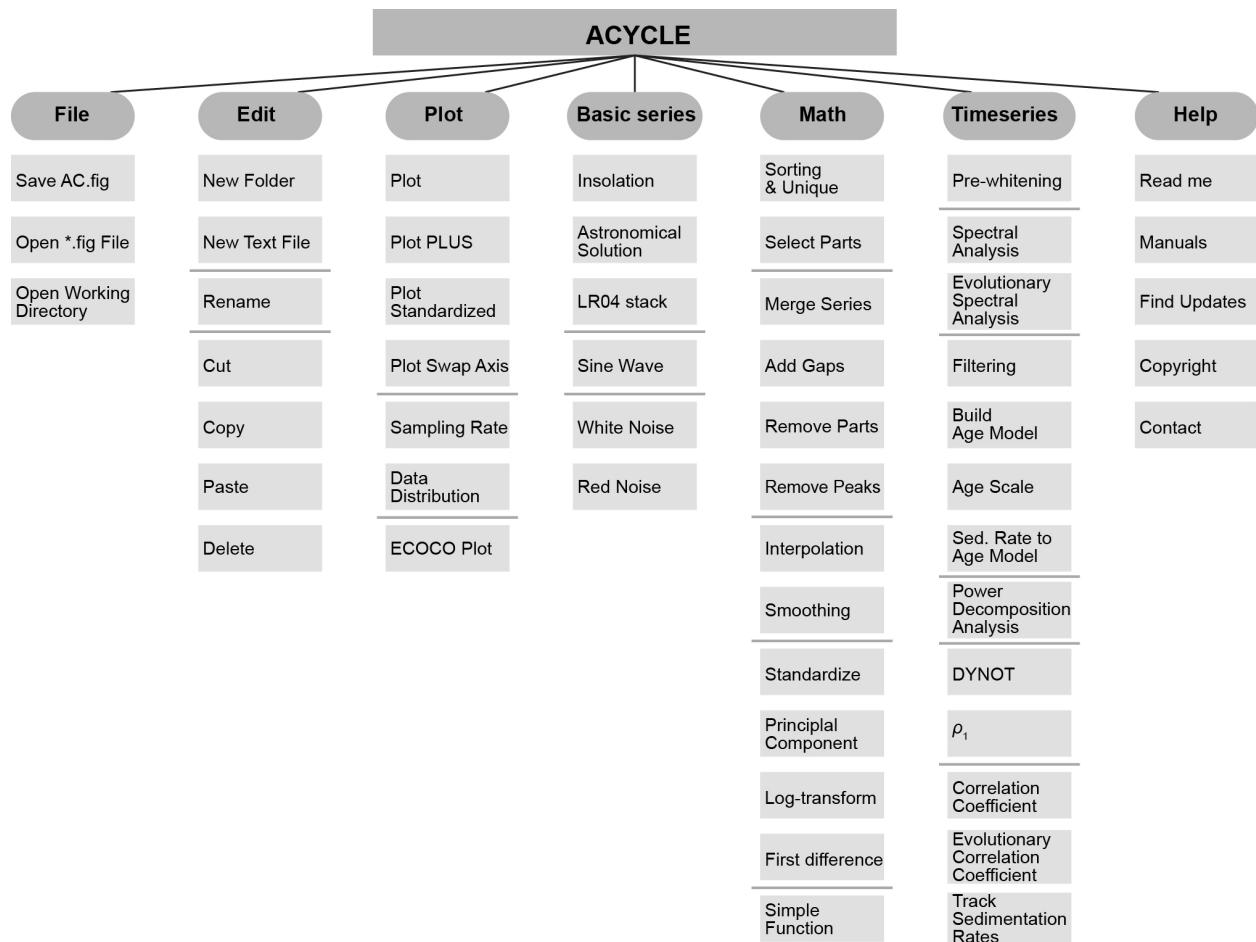
(Pre-whitening; Spectral Analysis; Evolutionary Spectral Analysis; Wavelet transform; Filtering; Amplitude Modulation; Build Age Model; Age Scale; Sedimentary Rate to Age Model; Power Decomposition Analysis; DYNOT; ρ_1 method; Correlation Coefficient; Evolutionary Correlation Coefficient; Track Sedimentation Rates)

Help

(Readme; Manuals; Find Updates; Copyright; Contact)



GUI of the *Acycle* program

The structure of the *Acycle* program

4.2 File

New Folder:

make a new empty folder with a user-defined folder name.

New Text File:

make a new empty *.txt file with a user-defined file name.

Save *.AC.fig file:

Save the current figure as an *.ac.fig file. This file enable users continue a suspended project.

For example, after running the eCOCO (evolutionary correlation coefficient), users may want to plot the eCOCO results anytime. One can save the current figure as an *.AC.fig file, then double click this *.AC.fig file and show “ECOCO plot” anytime.

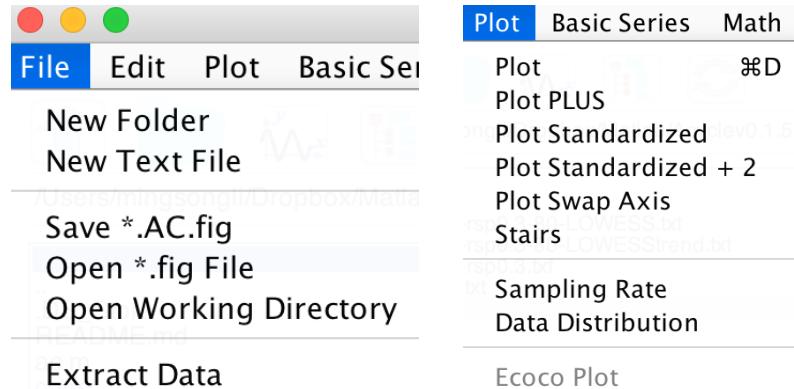
4.3 Edit

Rename:

Select one file, the “rename” function enable changing the name of the selected file.

Cut/Copy/Paste/Delete:

Just do it.



4.4 Plot

Plot:

A quick plot of the selected data file. Shortcut [Mac]: $\text{⌘} + \text{D}$; [Windows]: $\text{Ctrl} + \text{D}$

Plot PLUS:

An advanced plot of the selected data file (GUI below). One can change plot type, line, and marker styles, and control the axis.

Plot Standardized:

A quick plot of the standardized data file. Useful if one wants to compare 2 or more series.

Plot Standardized +2:

A quick plot of the standardized data file. Useful if one wants to compare 2 or more series.

Plot Swap Axis:

A quick plot, swap axis.

Stairs:

Stairs plot.

Sampling Rate:

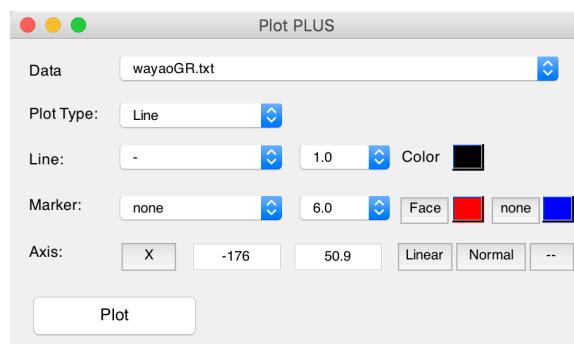
A quick plot showing the distribution of the 1st column (time/depth) of the selected data file.

Data Distribution:

A quick plot showing the distribution of the 2nd column (data) of the selected data file.

ECOCO Plot:

Plot eCOCO results from a *.AC.fig, or after running the eCOCO (Timeseries-eCOCO).

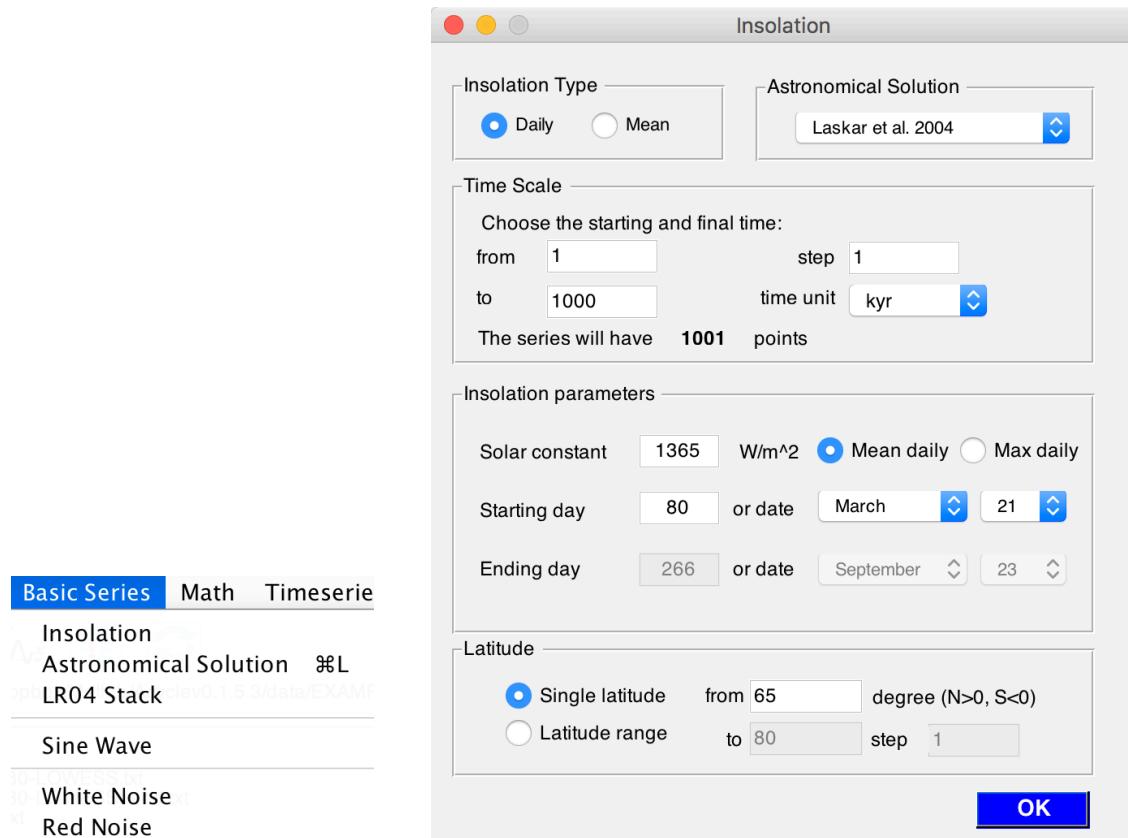


4.5 Basic Series

Insolation

A powerful GUI calculates the insolation using various astronomical solutions. Based on the MatLab code `insol.m` by Jonathan Levine (2001), UC Berkeley. This code was modified by Peter Huybers (Harvard) and edited by Mingsong Li (Penn State, 2018) for the *Acycle* software.

Only insolation series younger than 249,000 k.a. are available.

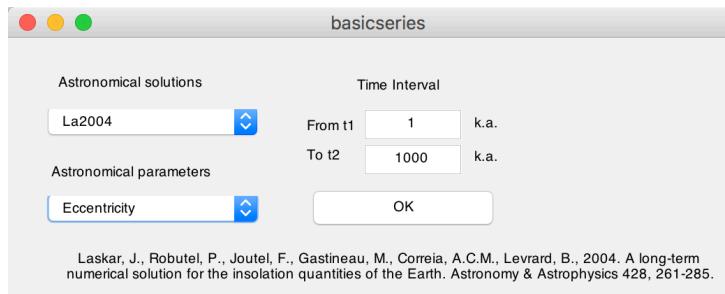


This GUI generates mean daily insolation series on March 21 for the past 1000 kyr (1-1000) at 65°N using the Laskar et al. (2004) solutions. The calculate uses a solar constant of 1365 w/m².

[NOTE: one has to click the Refresh button  in the main window to see the generated insolation series]

Astronomical Solution

A GUI generates astronomical solutions of Laskar et al. (2004); Laskar et al. (2011), and Zeebe (2017).

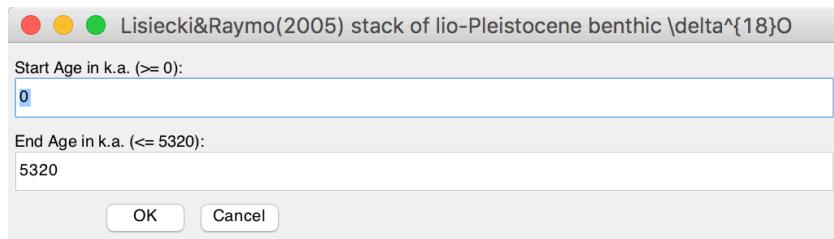


This GUI generates eccentricity series for the past 1 million years from 1 k.a. throughout 1000 k.a. using the La2004 solution ([Laskar et al., 2004](#)).

[NOTE: one has to click the Refresh button  in the main window to see the generated insolation series]

LR04 Stack

This function generates the classical LR04 stack of the Plio-Pleistocene benthic $\delta^{18}\text{O}$ record ([Lisiecki and Raymo, 2005](#)). The input time (below) should be within the interval of 0 and 5320 (k.a.).



This GUI generates LR04 stack from 0 to 5320 k.a.

Sine Wave

Generate a sine wave using user-defined parameters and the following equation:

$$Y = A * \sin(2\pi / T * X + Ph) + bias$$

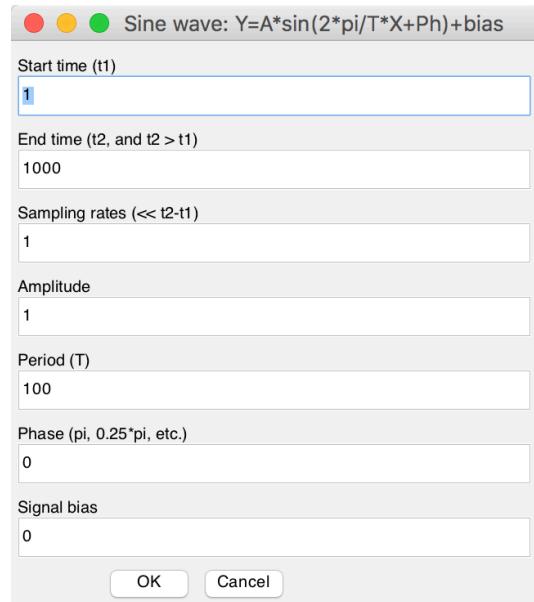
Where A is amplitude, T is period, X is a time series ranges from $t1$ to $t2$ and a sampling rate of dt, Ph is the phase, and bias is signal bias.

White Noise

This function generates the white noise.

Red Noise

This function generates the red noise using user-defined standard deviation, and autocorrelation coefficient (RHO-1, from 0 to 1).



This GUI generates a sine wave from 1 to 1000 unit with a sampling rate of 1 unit. Its amplitude is 1, with a period of 100 unit and zero phase shift and 0 signal bias.

4.6 Math

Sort/Unique/Delete-empty

This function will sort the selected data file like MS Excel's SORT function. If a dataset contains 2 or more data points with the same time/depth, then these data points will be replaced by their mean values.

New file name: *-su.txt or *-s.txt or *-u.txt

Select Parts

This function generates a new series from the selected data using user-defined 'start' and 'end' of the interval.

New file name: *-a-b.txt, where a is the "start" and b is the "end".

Merge Series

Two selected series may be merged if their first column are exactly the same.

New file name: mergedseries.txt.

Add Gaps

This function generates a new series based on the selected data file via adding a gap or gaps using user-defined location and duration of the gap(s). Format, comma delimited:

10.5, 3.2

Add a 3.2-unit gap at the depth/time of 10.5 unit, or

10.5, 3.2, 13.3, 1.5

Add a 3.2-unit gap at the depth/time of 10.5 unit and add the second 1.5-unit gap at the depth/time of 13.3 unit.

Remove Parts

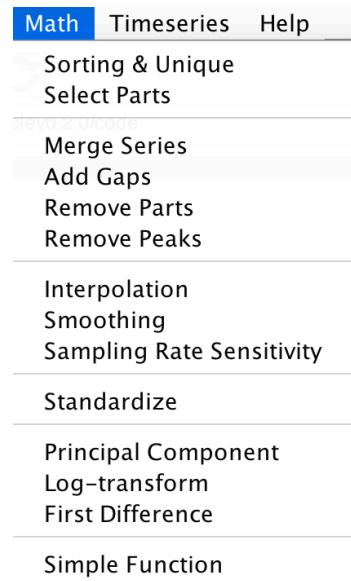
This function generates a new series based on the selected data file via removing an user-defined interval(s). Format, comma delimited

15, 3, 20.2, 4

Remove a 3-unit data at the 15 unit (remove 15-18-unit data), and remove the second interval of 20.2-24.2-unit.

Remove Peaks

This function generates a new series based on the selected data file via converting any (2nd column) data higher than the user-defined Maximum value to that value and any data smaller than Minimum value to that value.



Clipping

This function generates a new series based on the selected data file via clipping data higher or smaller than the user-defined threshold value.

Interpolation

Linear interpolation using MatLab's *interp1* function.

New file name: *-rsp0.3.txt, where 0.3 is user-defined interpolation sampling rate.

Smoothing:

Moving Average

This function generates a new series based on selected data file using n -points smoothing, where n is a user-defined parameter.

New file name: *-3ptsm.txt, means 3 points smoothing output.

Bootstrap

This function generates two new series based on selected data file using *user-defined* smoothing window, smoothing method, and number of bootstrap sampling.

New file name: *-WINDOW-METHOD-NUMBER-bootstp-meanstd.txt, mean and standard deviation data, and

*-WINDOW-METHOD-NUMBER-bootstp-percentile.txt, 0.5%, 2.275%, 15.865%, 50%, 84.135%, 97.725%, and 99.5% percentiles.

Gaussian Process

This function generates a new series based on selected data file using *user-defined* new x axis values using the Gaussian processes for machine learning (GPML) toolbox by Rasmussen (2010).

New file name: *-gp-meanstd.txt, mean and standard deviation data.

Standardize

Using MatLab's *zscore* function.

$Z = (X-u)/\sigma$, where X is the second column data, u is the mean of X , and σ is the standard deviation of X .

New file name: *-stand.txt

Principal Component

This function has different requirements of the data inputs. All column (including the first column) of data should be value, not depth or time.

Log-transform

This function generates a new data file based on selected data file using log10 transformation of the second column of the selected data.

$$X_i = \log_{10}(X_i)$$

New file name: *-log10.txt

First Difference

Differences using MatLab's diff function.

$Y = \text{diff}(X)$, calculates differences between adjacent elements of X .

New file name: *-1stdiff.txt

Derivative

Approximate derivatives (first, second, third, ...).

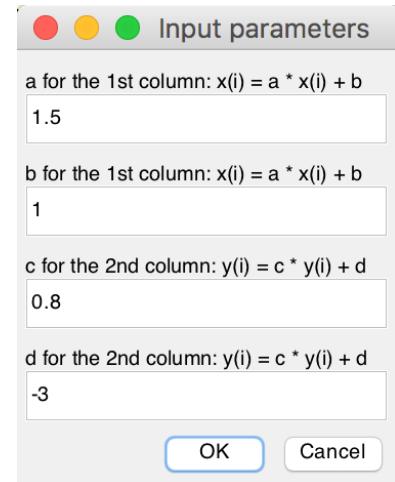
New file name: *-1derv.txt

Simple Function

This function is very useful. It generates a new data file based on the selected data file. Both columns (1st or X column and 2nd or Y column) can be modified. See below case study.

$$X_{(i)} = a * X_{(i)} + b$$

$$Y_{(i)} = c * Y_{(i)} + d$$



*The selected data: all value in the first column data will be transformed using the equation $X_{(i)} = 1.5 * X_{(i)} + 1$; and all value in the second column data will be transformed using the equation $Y_{(i)} = 0.8 * Y_{(i)} + (-3)$.*

New file name: *-new.txt

Utilities

Find max/min

Find max/min value within a user-defined interval. Output will be displayed in command window only.

Image:

Show Image

Plot selected image file.

RGB to Grayscale

Convert a image file in RGB format to a grayscale format, save new image .

New image name: *-gray.tif

Image Profile

Get the grayscale profile from a line constrained by two user-selected dots.

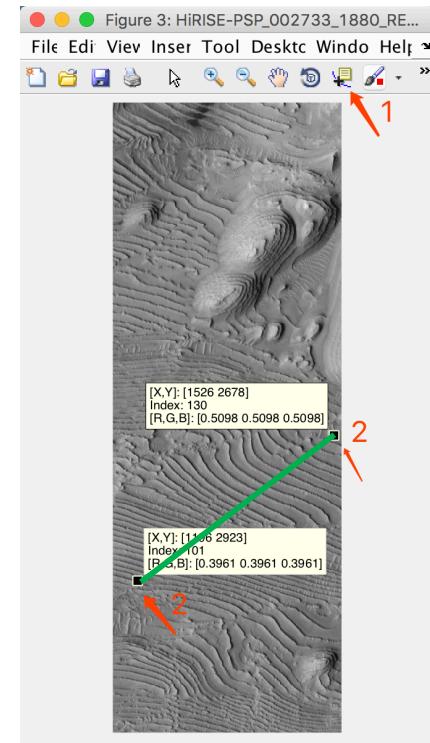
New file name: *-profile.txt %
grayscale profile

New file name: *-controlpoints.txt % location of two control points

Step 1: Choose the image file, select "Math - Image - Image Profile" function.

Step 2: Click data cursor tool (1), press ALT key and click 2 points.

Step 3: Press Enter key. Grayscale profile data will be picked up and saved along the green line.



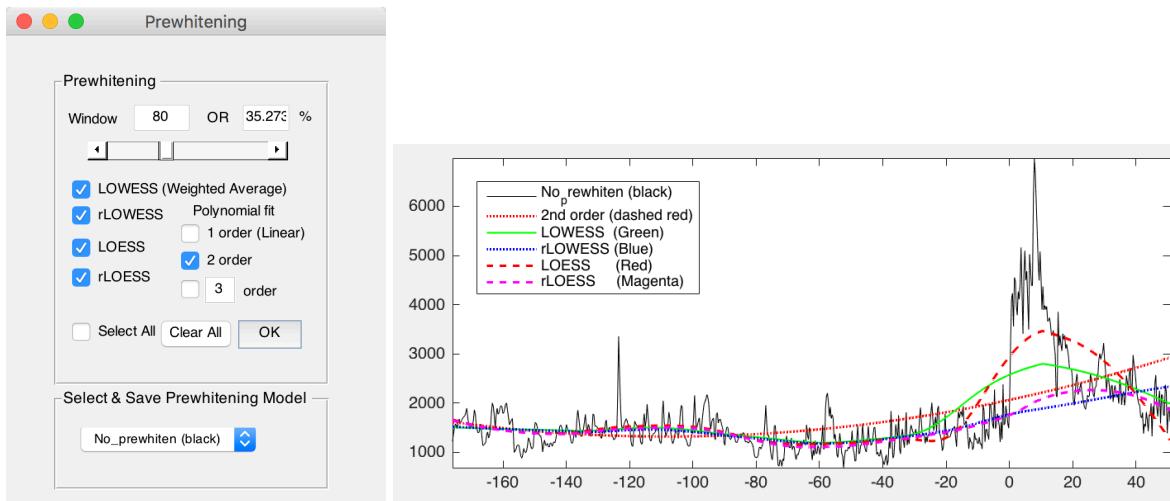
4.7 Time series

Prewhitenning

This prewhitening function generates 2 new data files based on the selected data file and user-defined parameters: window length and detrending method. Steps:

- (1) Select a data file in the Main Window
- (2) Select **Timeseries → Prewhitenning** menu
- (3) Type a window length or a percentage, or move the slider. Default value is 35% of the total length, that is, if a data length is 100 m, then a window is 35 m.
- (4) Tick one or more prewhitening method
- (5) Click **OK** button, wait for several seconds (up to a minute, depending on the length of the dataset and the speed of your machine). A new window will popup showing the data and its 35% trend(s).
- (6) In the ‘**Select & Save Prewhitenning Model**’ panel, select the preferred trend. The trend and detrended file will be displayed in the Main Window after user clicks the refresh button  in the Main Window.

New file names: *-80-LOWESS.txt AND *-80-LOWESStrend.txt



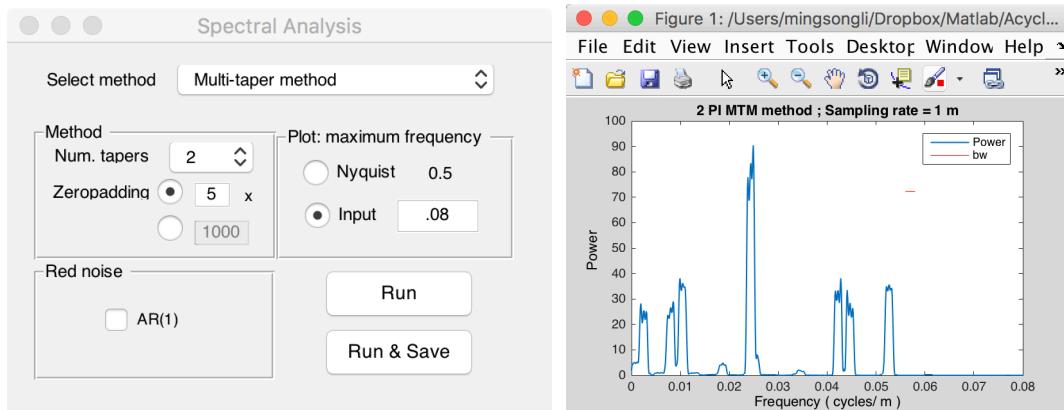
Spectral Analysis

This function conducts spectral analysis with user-defined parameters. Steps:

- (1) Select a data file in the Main Window
- (2) Select **Timeseries → Spectral Analysis** menu
- (3) Select one method for spectral analysis. Options are Multi-taper method (MTM) ([Thomson, 1982](#)), Lomb-Scargle spectrum ([Lomb, 1976](#); [Scargle, 1982](#)), and MatLab’s periodogram.
- (4) If Multi-taper method (MTM) is selected, then the Method panel may be changed. The default value is using 2π MTM, with a 5X zeropadding.
- (5) Plot panel: set the max frequency in the coming figure.

(6) Red Noise panel: AR(1) noise model using RedNoise.m by [Husson \(2014\)](#) and modified by Linda Hinnov.

(7) **Run** or **Run & Save** button, generates power spectrum (and save power spectrum data and AR(1) series)



2π MTM power spectrum of the La2004 astronomical solution (0-2 Ma)

New file name: *-2piMTM-CL.txt, means 2π MTM and confidence level series.

Evolutionary Spectral Analysis

This function conducts evolutionary spectral analysis with user-defined parameters. Steps:

(1) Select a data file in the Main Window.

Warning: The data file must be an evenly spaced depth/time series.

(2) Select **Timeseries → Evolutionary Spectral Analysis** menu

(3) Select Method. The default method is Fast Fourier transform (LAH) by Linda A. Hinnov ([Kodama and Hinnov, 2015](#)). Other options are MatLab's Fast Fourier transform and periodogram, multi-taper method (MTM) ([Thomson, 1982](#)), and Lomb-Scargle spectrum ([Lomb, 1976](#); [Scargle, 1982](#)).

(4) Input for evolutionary spectral analysis panel includes settings for plot frequencies. Default values from 0 to Nyquist ($f_{nyq} = 1 / (N * \Delta t)$), where N is the total number of data, and Δt is the sampling rate.

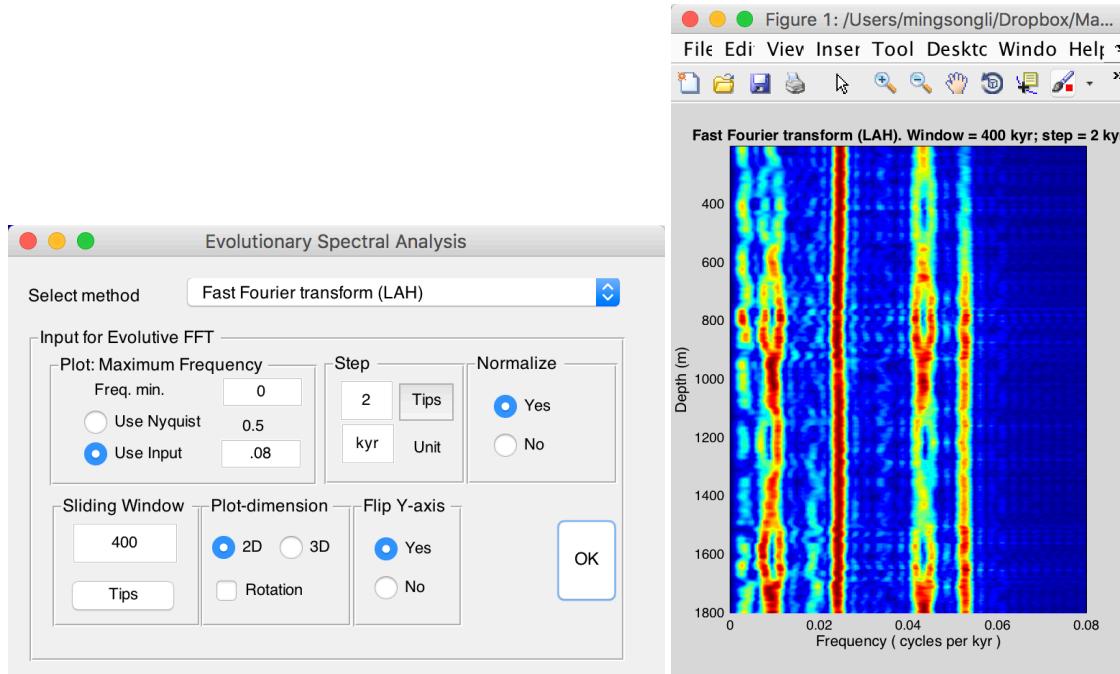
(5) Step of sliding windows. The default value should be sufficient for most paleoclimate projects. The unit may be *m*, *kyr*, etc.

(6) Sliding Window: **very important!** The length of the sliding window. The default value is 35% of the total length of the selected data. Tips: assuming the data series is dominated by 35 m cycles, the window may be 2-4 times of 35 m, that is, 70 to 140 m. A large window can smooth out the higher frequency signals while a small window cannot detect low-frequency signals.

(7) Plot-dimension: 2D or 3D with rotation option.

(8) Flip Y-axis: give me a try.

(9) **OK** button: generates a new figure showing the evolutionary spectral analysis results. No new files generated automatically.



Evolutionary FFT of the La2004 astronomical solutions using a 400 kyr sliding window and 2 kyr step

Wavelet transform

This wavelet analysis function conducts wavelet analysis (Torrence and Compo, 1998) with user-defined parameters. Steps:

(1) Select a data file in the Main Window.

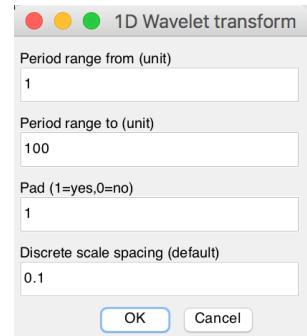
Warning: The data file must be an evenly spaced depth/time series.

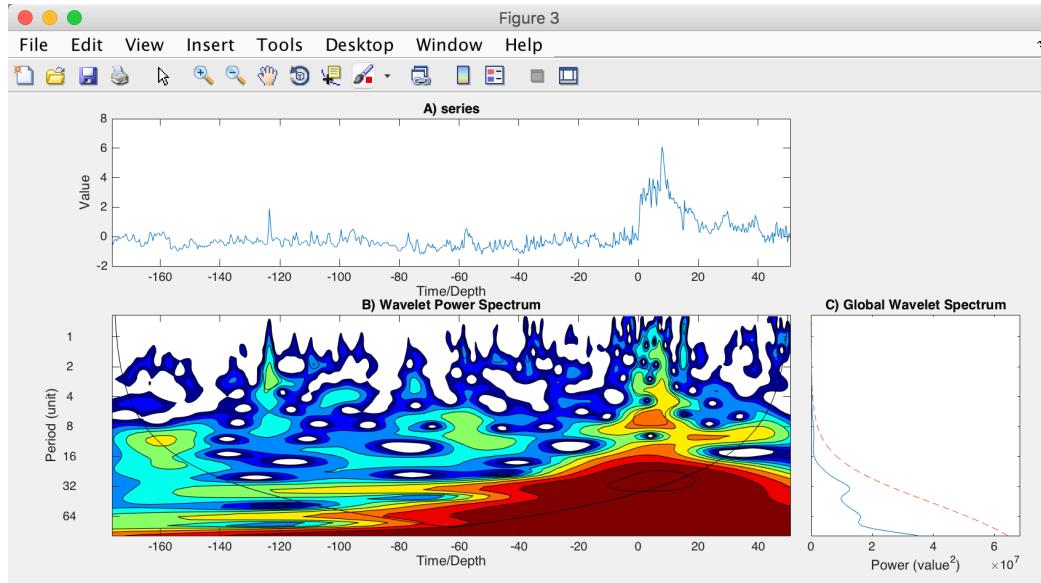
(2) Select **Timeseries → Wavelet Transform** menu

(3) Modify parameters

Period ranges from [the first line] to [the second line] unit.

Default values for all lines works well with the program. Users may need to modify the period range in the 2nd line using a smaller number (e.g., halved value).





Wavelet analysis of the Wayao gamma ray series. The series has been interpolated using a 0.3 m sampling rate.

Filtering

This function generates a filter output series based on the selected data file with user-defined parameters. Steps:

- (1) Select a data file in the Main Window.

Warning: The data file must be an evenly spaced depth/time series.

- (2) Select **Timeseries → Filtering** menu

(3) **Bandpass filter** panel: very important! Type min and center frequencies of the passband, the max frequency will be set automatically. The bandpass filters are MatLab's Butter, Cheby1, and Ellip filters and Gaussian, and Taner-Hilbert filters. The recommended filters are Gaussian filter and Taner-Hilbert filters code by Linda Hinnov ([Kodama and Hinnov, 2015](#)).

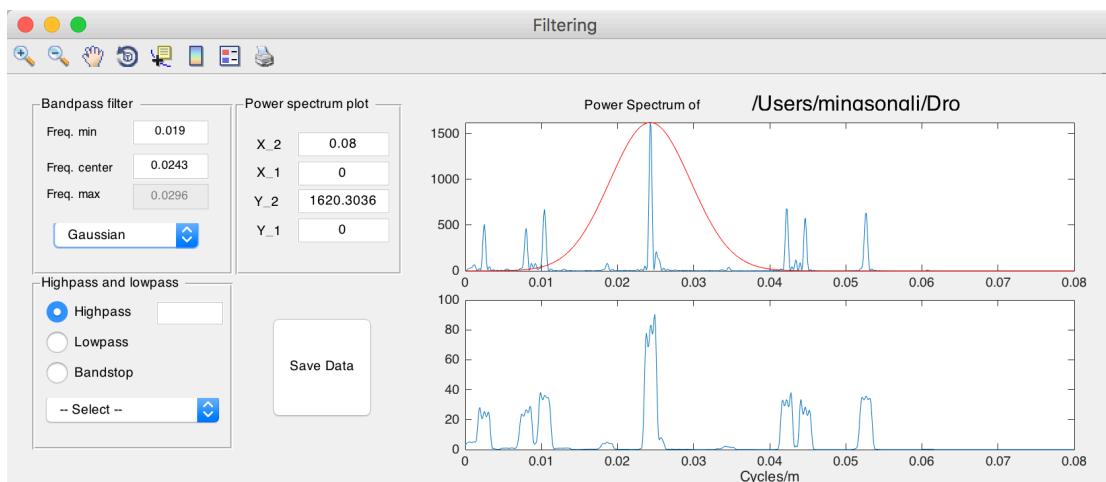
Tips: The Taner-Hilbert filter generates both filtered output series and the amplitude modulation of the filtered output series.

Click **Save Data** button, the filter outputs will be displayed after clicking the **refresh** button  in the Main Window.

(4) **Highpass and lowpass** panel: Two options are MatLab's Butter and Ellip filter. Type cutoff frequency in the text box and select a filter.

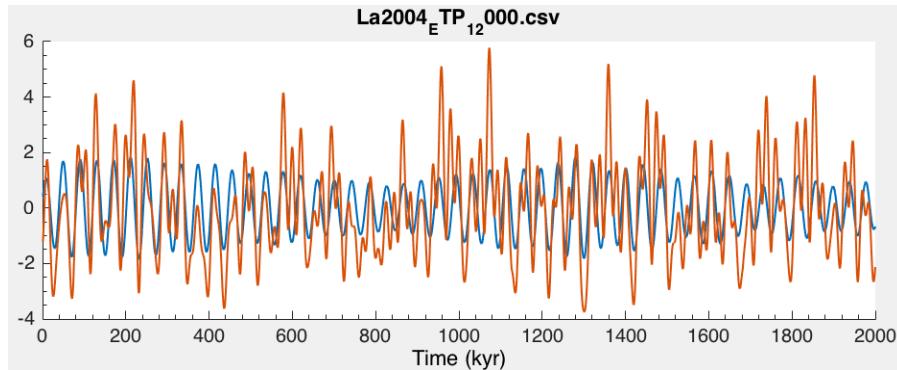
Click **Save Data** button, the filter outputs will be displayed after clicking the **refresh** button  in the Main Window.

- (5) Power spectrum plot: give options for display the power spectrum in the right of the GUI.



New file name: *-gaus-0.0243+-0.0053.csv, means filtered output series using gauss filter and a 0.0243 ± 0.0053 cycles/unit bandpass.

*-Tan-0.03+-0007.csv and *-Tan-0.03+-0007-AM.csv, mean filtered output series using Taner-Hilbert filter and a 0.03 ± 0.007 cycles/unit bandpass, with its amplitude modulation file saved.



Plot original La2004 solutions and filtered 41 kyr cycles

Build Age Model

This function generates an age model file from a filter output data file. Steps:

(1) Assuming the filtering wavelength generates a filtered 35 m cycle series. The 35 m cycles are assumed to be 405 kyr long eccentricity cycles. This filtered data file should be selected.

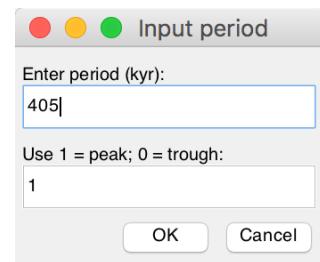
(2) Select **Timeseries → Build Age Model** menu

(3) In the pop-up window, enter 405 and 1, and click OK button.

This generates a new age model series via assigning every peak of 35 m cycles as peaks of the 405 kyr cycles.

New file name: *-agemode1-405-max.csv,

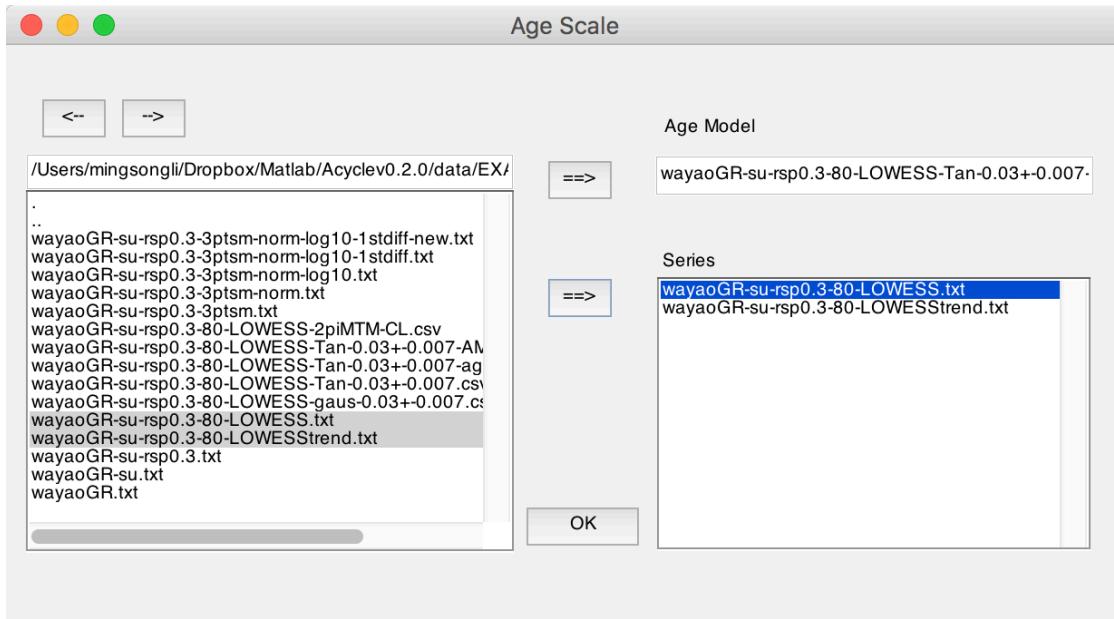
means an age model file using filtered wavelength peaks as 405 kyr anchors.



Age Scale

This function conducts depth-to-time transformation in a new standalone GUI. Steps:

- (1) Select 1 (ONE) age model file, click the top \Rightarrow button to record this file as an age model file.
- (2) Select 1 or more data files, click the bottom \Rightarrow button to record this file (these files) as series needs to be transformed.
- (3) Click the **OK** button. The transformed series can be displayed after clicking the Refresh button  in the Main Window.



New file name(s): *-TD-name-of-agemodel-file.csv

(Tips) Change directory using \leftarrow or \rightarrow button

Sedimentary Rate to Age Model

Assuming you want to generate an age model file from a sedimentary rates file (2 columns: depth and sedimentation rate), this function generates the age model working well with acycle software.

Power Decomposition Analysis

This function subtracts power/variance within a user-defined frequency band. The code written by Mingsong Li and Linda Hinnov has been published in [Li et al. \(2016\)](#). Time-dependent amplitude modulations in the obliquity component were obtained from 2π multi-taper variance (power) spectra calculated along a sliding time window ([Meyers et al., 2012](#)) using the Matlab script *pda.m* (also available at <https://doi.pangaea.de/10.1594/PANGAEA.859147>). Steps:

- (1) Select the original data file.

Warning: The data must be evenly spaced data in the first column. And the unit must be in kyr.

(1) Type paired frequency bands; space delimited. If a dominated frequency is 1/33, then a 1/45 1/25 frequency band is used

(2) Sliding window in kyr, a 500 kyr is used in [Li et al. \(2016\)](#)

(3) Time-bandwidth product, ‘2’ means 2π MTM method will be used.

(4) cutoff frequencies, min = 0, max should cover all Milankovitch frequencies.

DYNOT

Dynamic noise after orbital tuning. Detect non-orbital variances from a tuned series. See **Chapter 5. DYNOT model Description**. See [Li et al. \(2018a\)](#) for details about this method.

ρ_1 method

This function conducts either single run or Monte Carlo simulations of lag-1 autocorrelation coefficient (ρ_1) using a sliding window. It works with both depth series and time series.

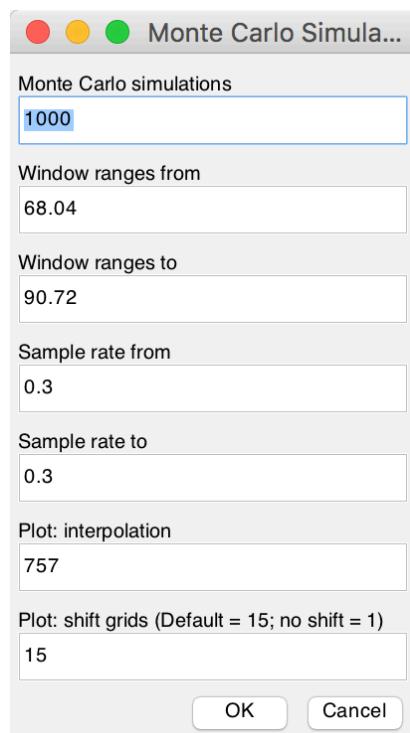
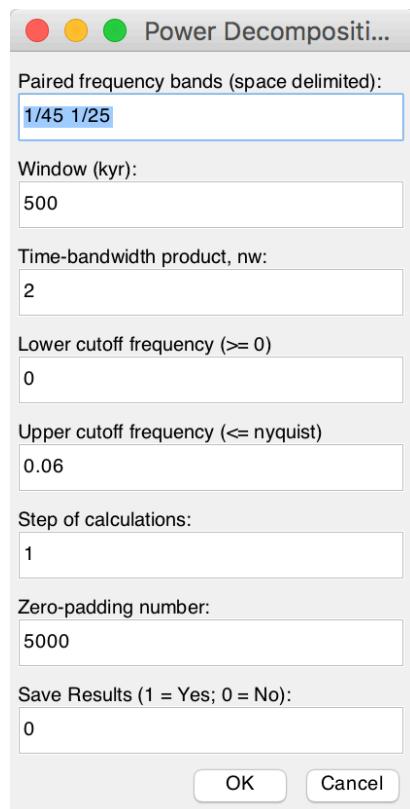
The “**Single run**” requires the input of “window” and “interpolation sampling rate”.

The “**Monte Carlo**” requires several parameters: Number of Monte Carlo simulations (default is 1000), sliding window ranges from *win1* to *win2*, and a sampling rates from *sr1* to *sr2*, and plot settings (interpolation and shift grid).

See [Li et al. \(2018a\)](#) for details about the parameters and significance of this method.

Correlation Coefficient (COCO)

This function addresses two fundamental issues in cyclostratigraphy and paleoclimatology: identification of astronomical forcing in sequences of stratigraphic cycles, and accurate evaluation of sedimentation rates. The technique considers these issues part of an inverse problem and estimates the product-moment correlation coefficient between the power spectra of astronomical solutions and paleoclimate proxy series across a range of test sedimentation rates. The number of contributing astronomical parameters in the estimate is also considered. Our estimation procedure tests the hypothesis that astronomical forcing had a significant impact on proxy



records. The null hypothesis of no astronomical forcing is evaluated using a Monte Carlo simulation approach. Details are included in ([Li et al., 2018b](#)).

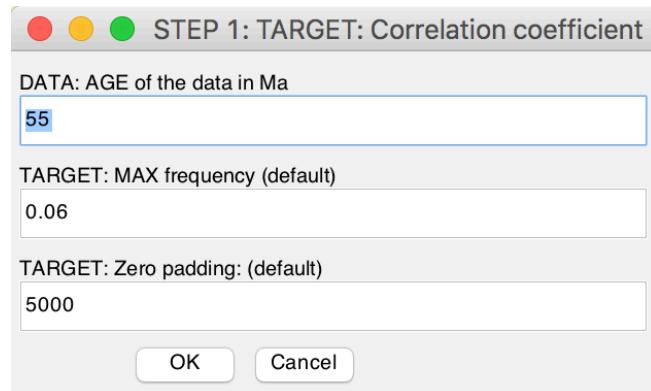
Step 1: settings for generating target power spectrum

Select a depth series (interpolated, detrended), select **Timeseries --> Correlation Coefficient** menu

Note: the data series must have a unit in meter.

Type the approximate age for the depth series, the unit is million years ago (Ma).

Target frequency ranges from 0 cycle/kyr to the given “MAX frequency”. Default values are recommended for the depth series with age less than 250 Ma. For the depth series older than 250 Ma, the MAX frequency may be set to 0.08. This is because the precession cycle can be very short than 16 kyr.



Step 2: astronomical solution [optional]

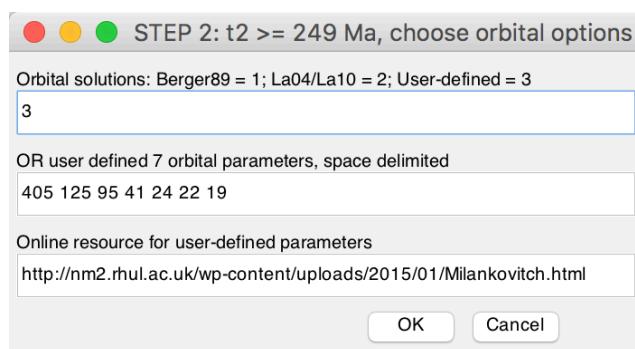
If the age of the data in Ma is larger than 249 Ma, users need to select which astronomical solution should be used.

1 = Berger89 solution ([Berger et al., 1989](#)),

2 = Laskar 2004 solution ([Laskar et al., 2004](#)),

3 = user-defined solution and the second box should be filled by 7 astronomical periods.

Online resource for user-defined astronomical parameters may be found at
<http://nm2.rhul.ac.uk/wp-content/uploads/2015/01/Milankovitch.html> ([Waltham, 2015](#)).



Step 3: settings for generating data power spectrum

MIN sedimentation rate (cm/kyr):

MAX sedimentation rate (cm/kyr):

STEP sedimentation rate (cm/kyr): tested sedimentation rates range from *MIN* to *MAX*, with a step of *STEP* cm/kyr. In the following example, the tested sed. rates are 1, 1.5, 2, 2.5, 3, ..., 29.5, and 30 cm/kyr.

Number of simulations: 200-600 simulations are suggested for an initial run. And 2000 simulations generate publication quality results.

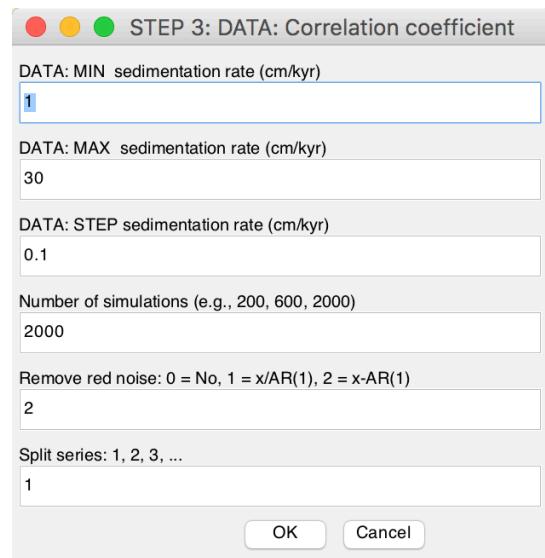
Remove red noise: 0 = no removing; else removing red noise:

1 = power spectrum / AR(1) series and those less than AR(1) series are set to 0;

2 = power spectrum - AR(1) series and those less than 0 are set to 0.

Split series: 1, 2, 3. The series is split into 2 or more slices.

Click the **OK** button, Monte Carlo simulation steps can be displayed in the Command Window of MatLab. A log file will be generated recording all parameters used in the correlation coefficient analysis. The user needs to decide which figure output should be saved or not.



Evolutionary Correlation Coefficient (eCOCO)

The method is applied using a sliding stratigraphic window to track variable sedimentation rates along the proxy series, in a procedure termed “eCOCO” (evolutionary correlation coefficient) analysis. ([Li et al., 2018b](#))

Waning: the data series must have a unit in meter.

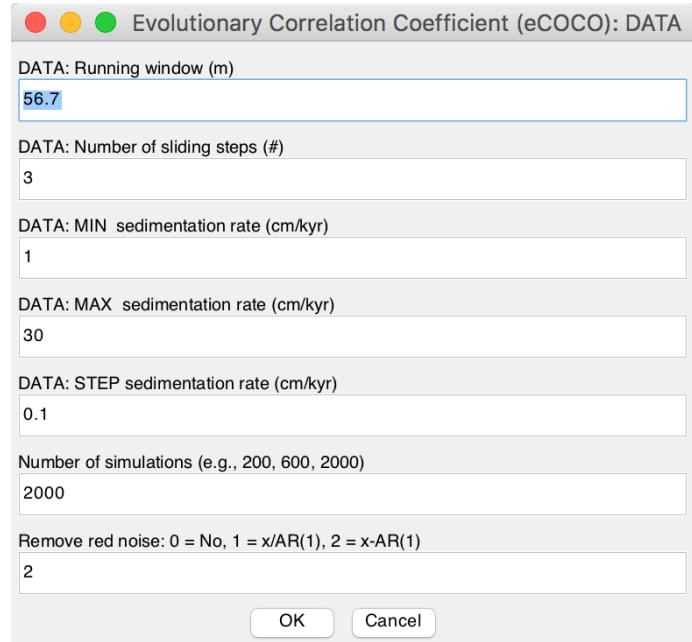
Step 1: same as that in COCO.

Step 2: same as that in COCO.

Step 3: most parameters are the same as those in COCO (see above). Two new parameters:

DATA: running window (m): default window is 35% of the total length of the data series.

DATA: Number of steps (#): sliding steps. The default value will give about ~300 sliding windows for publication quality results.



Click the **OK** button, Monte Carlo simulation steps can be displayed in the Command Window of MatLab. A log file and the related *.AC.fig file will be generated recording all parameters used in the

evolutionary correlation coefficient analysis. The user needs to decide which figure output should be saved or not.

Tips: Users may save the main window using “File” → “save ac.fig” menu anytime. This will save the data stored in the main window figure, and the user doesn’t have to re-run the eCOCO using the same parameters.

Tips: User can plot eCOCO results anytime at “Plot” → “ECOCO plot” menu.

Q: Which window should I use?

A: Well, if your series is dominated by 35 m cycles (405 kyr), then a 70 m window (= 35 *2) may be good to keep the balance: A large window eCOCO losses resolution of variable sedimentation rates, and a small window may not give correct results.

Track Sedimentation Rates

Not finish yet...

4.8 Help

Readme

Update log

Manuals

Open the folder of user’s guide

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www.mingsongli.com/acycle

<https://github.com/mingsongli/acycle>

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5. DYNOT model Description

[Li et al. \(2018a\)](#) developed a dynamic noise after orbital tuning, or DYNOT model for the recognition of sea-level variations based on the dynamic non-orbital signal in climate proxy records after subtracting orbital, i.e., astronomically forced climate signal. The DYNOT model is supplemented by a second, independent lag-1 autocorrelation coefficient, or ρ_1 model, which forms the basis of a well-established statistical method for red noise estimation of time series. DYNOT and ρ_1 modeling of a GR series of ODP Site 1119 over the past 1.4 myr correlates with the classic low-passed $\delta^{18}\text{O}$ sea-level curve, demonstrating the efficacy of the sedimentary noise model.

5.1 Startup

* Timeseries - DYNOT

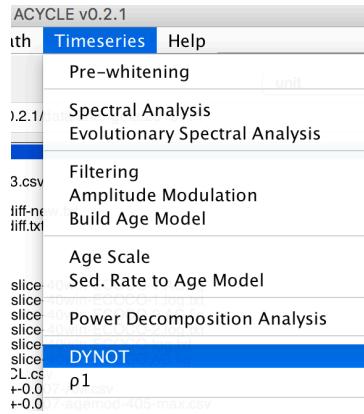


Fig. 1. MatLab workspace for the DYNOT model.

The DYNOT sea-level model GUI is as follows (Fig. 2):

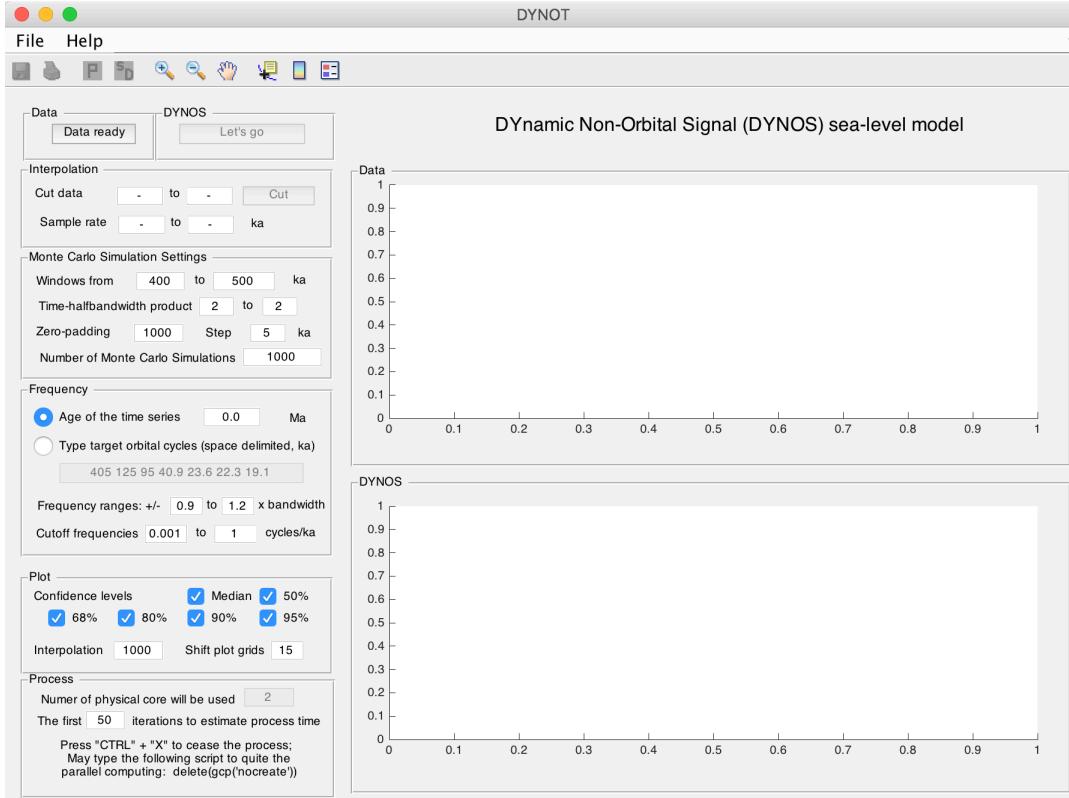


Fig. 2. The DYNOT model

5.2 Import Data

*data for the DYNOT model (support data in *.csv and *.txt format)*

Name:	data		
Length:	$m \times 2$	% must be a 2-column dataset	
Column 1:	time	% unit must be in ka;	
Column 2:	value		

Notes:

#1: Proxy data is assumed to be sensitive to water-depth related noise at your section/core.

#2: There is no requirement for interpolation, normalization, or removing long-term trend (i.e., pre-whitening) of the dataset.

#3: Extreme values should be removed.

#4: Both increasing-upward and decreasing-upward time series are valid.

from *.txt or *.csv file

In the DYNOT menu: Select “File” → “Import Data (*.txt, *.csv)” → Select data (choose “1119_gr_1400de_finetuned.txt” or “1119_gr_1400de_finetuned.csv”) → Click “Open”

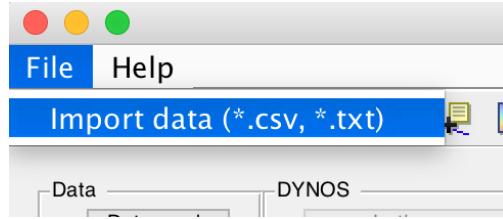


Fig. 3. Load data to DYNOT model.

5.3 Settings

Yellow: load data and run the model.

Red: Key settings. Check before running the model.

Green: Optional settings. Default values are okay for most running.

5.3.0. Click on **Data ready** (button) to load data into the DYNOT model.

5.3.1. Cut data (*optional*): These settings automatically show the beginning and the end of the time series, i.e., time span of dataset. Unit is ka. If you want to choose a different interval, just type two new ages and click **Cut** button.

5.3.2. Sampling rates (*optional*): These show a range of sample rates covering 90% of sample rates (Green Box 20 in Fig. 4). Unit is ka. A Monte Carlo method of hypothesis testing and the multi-taper method (MTM) of power spectral analysis are to be undertaken, and so resampling must be applied. Sampling rates of proxy datasets in time are always greater than zero and so are non-normally distributed. Therefore, the Weibull distribution is used to represent sampling rate distributions for uncertainty analysis in the DYNOT model. To avoid an ultra-low or ultra-high, unrealistic sampling rate created by the Weibull distribution algorithm, we set the 5th and 95th percentiles of sampling rates of the data as default, lower and upper limits of the generated, Weibull-distributed sampling rates.

5.3.3. Windows: These values set sliding window range. Moving window length in units of time (<< total data length). Unit is ka.

Different windows in the DYNOT model can affect results in two ways.

- (1) The DYNOT model with a large window will shorten DYNOT results, and the model with a small window will generate longer DYNOT results, $N_r = N_{data} - window + 1$, where N_r is total number of DYNOT values of each simulation, N_{data} is total number of interpolated data points, and $window$ is the running window employed.
- (2) The DYNOT model with a small running window generates higher resolution results, however, the variance of low-frequency cycles and total variance diminish simultaneously, which leads to increased uncertainty in non-orbital signal ratio estimation.

The DYNOT model with a small running window also increases the MTM power spectrum bandwidth (i.e., reduces frequency resolution). The expected sea-level variations of interest in the Early Triassic are 10^4 to 10^6 year-scale, i.e., the fifth to

third-order sequences, therefore a comparable or shorter time window (e.g., 300-500 kyr, 400 kyr or shorter) should be adopted for DYNOT modeling.

- 5.3.4. Time-bandwidth product (*optional*): Time-bandwidth product of discrete prolate spheroidal sequences used for window. Typical choices are 2, 5/2, 3, 7/2, 4.
- 5.3.5. Zero-padding (*optional*). zero-padding number, e.g., 1000.
- 5.3.6. Step (*optional*). step of calculations; default is 5 ka.
- 5.3.7. **Number of Monte Carlo Simulations:** default is 1000. Maybe use 100 or 300 for a trial running. Recommended value for publication is >5000.
- 5.3.8. **Age of the time series:** The age in Ma will be used to estimated target orbital cycles in 5.3.9. You can use either 5.3.8 or 5.3.9 to tell the DYNOT model the target cycles.
- 5.3.9. **Target orbital cycles** (space delimited, in ka): 6 orbital cycles of long-eccentricity (405), short-eccentricity (125 and 95), obliquity (40.9 or shorter), precession (23.6, 22.3, and 19.1 or shorter). This is age dependent (see 7.8). The 405, 125, and 95 kyr cycles are assumed to be invariant through time. While the obliquity = $41 - 0.0332 \times \text{age}$; precession 1 = $23.75 - 0.0121 \times \text{age}$; precession 2 = $22.43 - 0.0121 \times \text{age}$; precession 3 = $19.18 - 0.0079 \times \text{age}$. These calculations are from [Yao et al. \(2015\)](#), and are based on the La2004 astronomical model ([Laskar et al., 2004](#)).

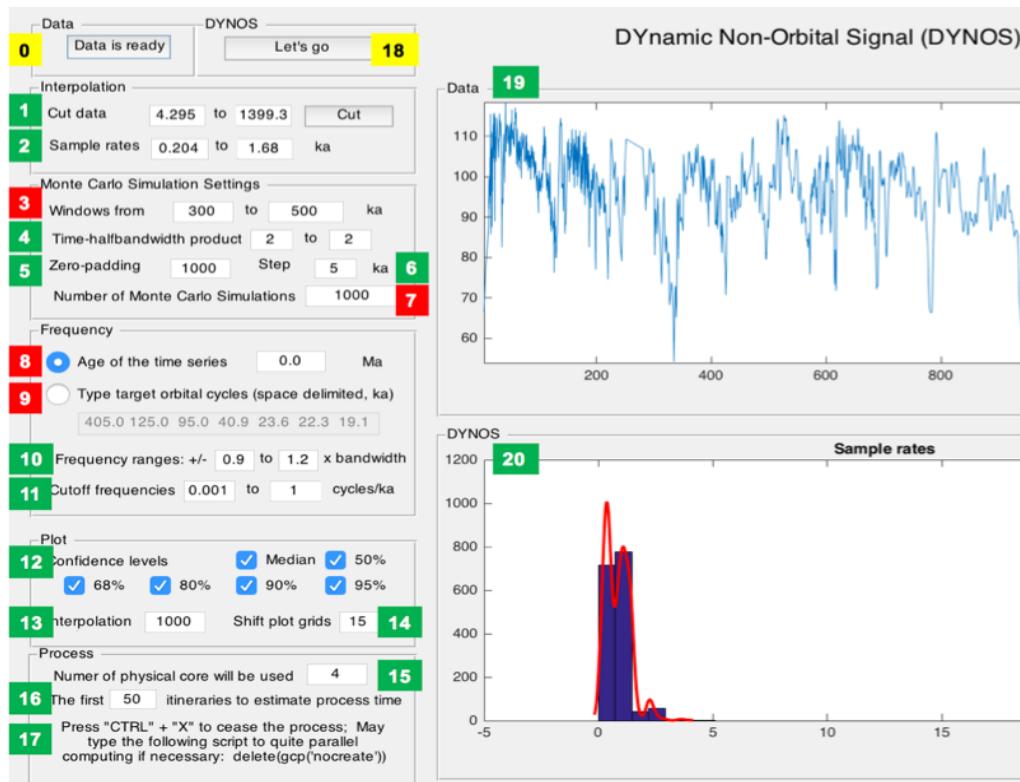


Fig. 4. Settings of the DYNOT model.

Yellow: load data and run the model.

Red: Key settings. Check before running the model.

Green: Optional settings. Default values are okay for most running.

- 5.3.10.** Frequency ranges (*optional*): For the definition of the non-orbital signal ratio by [Li et al. \(2018a\)](#), cutoff frequencies and their bandwidths are crucial for estimation of variances of eccentricity, obliquity and precession signals. We vary each cutoff frequency assuming a uniform distribution with cutoff frequency ranges at $\pm 90\%$ to $\pm 120\%$ bandwidth. Here the bandwidth (*bw*) equals $nw/window$, where *nw* is time-bandwidth product of discrete prolate spheroidal sequences, and *window* is the running window.
- 5.3.11.** Cutoff frequencies (*optional*): lower cutoff frequency (> 0) for estimation of total variance and upper cutoff frequency ($<$ Nyquist frequency) for estimation of total variance.
- 5.3.12.** Confidence levels (*optional*): default values show median and confidence levels (e.g., 50%, 68%, 80%, 90%, and 95%) of the DYNOT results.
- 5.3.13.** Interpolation (*optional*): In 5.3.3, a smaller *Nr* compared to *Ndata* leads to a “no data” effect at the very beginning and/or very end of the DYNOT results. To avoid this problem and to provide a better constraint for noise estimation, technically, the DYNOT model is interpolated and randomly shifts and plots simulation results of a single iteration at the same time scale of the dataset, although the plots also generate relatively smoothed DYNOT spectra when a gap is shorter than $2 \times window$. Here 1000 is adequate for the DYNOT model.
- 5.3.14** Shift plot grids (*optional*): See 5.3.13 for interpretation. Default is 15. One can also use 15-30 for the better shape of the beginning and the end of the DYNOT spectra.
- 5.3.15.** Number of physical cores (*optional*): This detects the physical cores of the CPU of the computer.
- 5.3.16.** Number of itineraries to estimate the process time (*optional*): To estimate process time of the time-consuming DYNOT model, the model will run some itineraries. Default is 50.
- 5.3.17.** Emergency note: Press “Ctrl” + “C” to cease the DYNOT process before the parallel computing. Press “Ctrl” + “X” to cease the DYNOT process during the parallel computing. You may need to type the following script in the command window to quite parallel computing.
- ```
>> delete(gcp('nocreate'))
```
- 5.3.18.** Click the button to run the model.
- 5.3.19.** A window shows the dataset.
- 5.3.20.** A window shows sample rates of the dataset OR the DYNOT spectrum of the dataset.

## 5.4. Running the DYNOT model

Click the **Let's go** button to run the DYNOT code. In the command window, the estimated running time will appear:

```
16:21:20 Begin the process ...
```

```

16:22:54 First 50 iterations suggest: remain >= 0h:7m:27sec
% The model runs the first 50 iterations to estimate that the total running time
will last ca. 7 minutes 27 seconds. The real run-time may be 10s seconds to
several minutes longer than this estimate.
Starting parallel pool (parpool) using the 'local' profile ...
connected to 4 workers.
16:23:07 Current iteration takes 1.11 seconds
16:23:08 Current iteration takes 1.21 seconds
16:23:15 Current iteration takes 1.19 seconds
16:26:26 Current iteration takes 1.38 seconds
% Start parallel computing and show time of each iteration.
Parallel pool using the 'local' profile is shutting down.
>> Done. % Stop parallel computing and display the DYNOT result (Fig. 5).

```

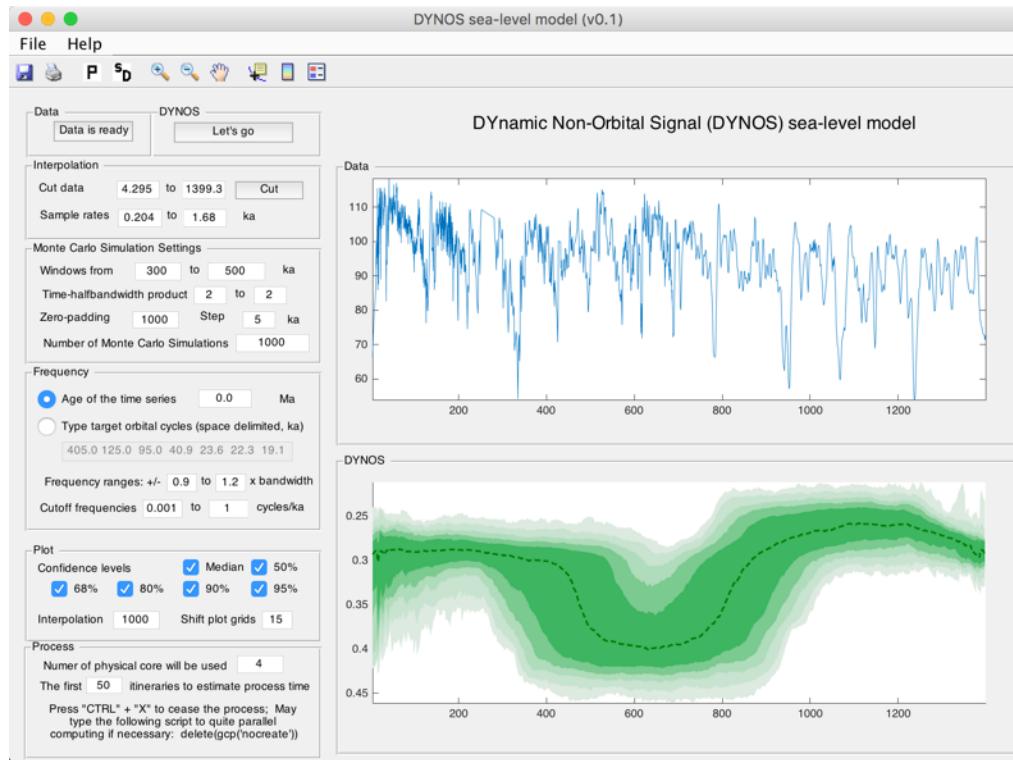


Fig. 5. DYNOT sea-level model of the gamma-ray series at ODP site 1119 from 0 to 1.4 Ma.

## 5.5. Output Files

After running the DYNOT model, the GUI menu (Fig. 6) can be used to:

- #1: save a MatLab-fig in the working directory entitled “plots\_.fig”.
- #2: save a PDF file of the plots in the working directory entitled “plots\_.pdf”
- #3: pop-up display the DYNOT spectrum in a new window.
- #4: save DYNOT output data in the working directory entitled “result\_handles.mat”.

**Caution: Change names of output files, or they will be overwritten by new files.**

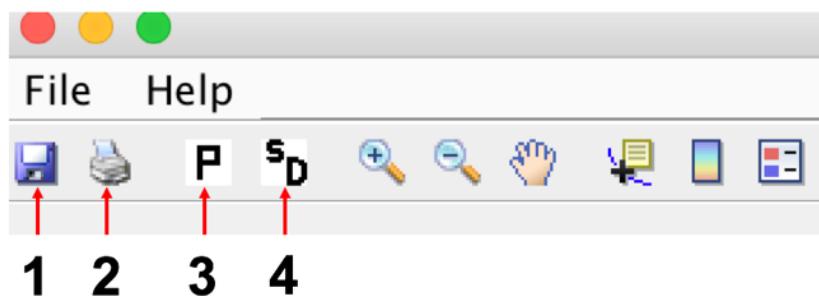


Fig. 6. Output files

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