



The High-redshift Clusters Occupied by Bent Radio AGN (COBRA) Survey: Investigating the Role of Environment on Bent Radio AGNs Using LOFAR

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Abstract

Bent radio active galactic nucleus (AGN) morphology depends on the density of the surrounding gas. However, bent sources are found inside and outside clusters, raising the question of how environment impacts bent AGN morphology. We analyze new LOw-Frequency Array the LOFAR Two-metre Sky Survey (LoTSS) Data Release II observations of 20 bent AGNs in clusters and 15 not in clusters from the high- z Clusters Occupied by Bent Radio AGN (COBRA) survey ($0.35 < z < 2.35$). We measure the impact of environment on size, lobe symmetry, and radio luminosity. We find that the most asymmetric radio lobes lie outside of clusters, and we uncover a tentative anticorrelation between the total projected physical area and cluster overdensity. Additionally, we, for the first time, present spectral index measurements of a large sample of high- z bent sources using LoTSS and Very Large Array Faint Images of the Radio Sky at Twenty-centimeters (VLA FIRST) observations. We find that the median spectral index for the cluster sample is -0.76 ± 0.01 , while the median spectral index for the noncluster sample is -0.81 ± 0.02 . Furthermore, 13 of 20 cluster bent AGNs have flat cores ($\alpha \geq -0.6$) compared to 4 of 15 of noncluster bent AGNs, indicating a key environmental signature. Beyond core spectral index, bent AGNs inside and outside clusters are remarkably similar. We conclude that the noncluster sample may be more representative of bent AGNs at large offsets from the cluster center (> 1.2 Mpc) or bent AGNs in weaker groups rather than the field.

Unified Astronomy Thesaurus concepts: High-redshift galaxy clusters (2007); Radio active galactic nuclei (2134)

1. Introduction

Galaxy clusters are the largest gravitationally bound structures in the Universe. Since galaxy clusters host large populations of similarly aged galaxies, clusters offer a unique lens to study galaxy evolution. Such studies require a large sample of similar clusters across a range of redshifts. While there are many well-studied low- z clusters that have been observed across the entire electromagnetic spectrum (e.g., Eisenhardt et al. 2007), as we probe higher redshifts, the number of well-studied clusters dwindles. However, large optical and infrared surveys performed with telescopes like the Spitzer Space Telescope and the Wide-field Infrared Survey Explorer (WISE) have led to a recent boom in the number of confirmed high- z clusters and cluster candidates (e.g., Muzzin et al. 2009; Wilson et al. 2009; Wylezalek et al. 2013; Gonzalez et al. 2015; Paterno-Mahler et al. 2017; Gonzalez et al. 2019). Recently, the Massive and Distant Clusters of WISE Survey (MaDCoWS) identified thousands of cluster candidates at $z \approx 1$ (e.g., Gonzalez et al. 2019). While surveys like MaDCoWS are excellent cluster finders, they preferentially identify massive, Coma cluster-like progenitors. However, as

shown in low- z cluster surveys (e.g., Miller et al. 2005; Rykoff et al. 2014), there are far more low-mass clusters and groups. As such, with the onset of large data sets and new, higher-resolution observations coming from JWST, LSST, and DESI, we find a conundrum in how to best identify potential lower-mass clusters at high- z where projection effects and large observational foreground and background contamination can cause false detections.

One strong tracer for finding high- z galaxy clusters and their progenitors are radio active galactic nuclei (AGNs; e.g., Minkowski 1960; Blanton et al. 2003; Wylezalek et al. 2013, 2014; Cooke et al. 2015, 2016; Noiro et al. 2016, 2018; Moravec et al. 2019, 2020; Shen et al. 2021). Specifically, the Clusters Around Radio-Loud AGN (CARLA) survey found that $\approx 92\%$ of their sample of radio-loud AGNs were in overdense systems and that 55% were in clusters (Wylezalek et al. 2013). Based on the density distribution of massive clusters/protoclusters and radio-loud AGNs, Hatch et al. (2014) found that approximately all protoclusters host radio-loud AGNs at $1.3 < z < 3.2$, making them an ideal tracer. Importantly, radio AGNs are also found in clusters across a range of masses and cluster dynamical/evolutionary states (e.g., Blanton et al. 2001; Wing & Blanton 2011; Cooke et al. 2015; Cai et al. 2017; Garon et al. 2019).

Despite radio AGNs being a very strong tracer, not every radio AGN is in a rich cluster (e.g., Blanton et al. 2000;



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Wing & Blanton 2011; Wylezalek et al. 2013; Paterno-Mahler et al. 2017). To better trace large-scale cluster environments with radio AGNs, the morphology of the AGNs should be considered. The standard classification scheme of radio AGNs is the Fanaroff-Riley (FR) system (Fanaroff & Riley 1974), which divides radio AGNs into FRI and FRII sources based on their radio luminosity ($L_{1.44\text{ GHz}} \approx 10^{25}\text{ WHz}^{-1}$) and the location of the brightest portion of the core and lobes. FRIs have brighter cores and lobes that fade at the edges, while FRIIs have large bright lobes farther from the AGN core (Fanaroff & Riley 1974). FRIs are commonly in clusters at both low and high redshift (e.g., Longair & Seldner 1979; Prestage & Peacock 1988; Hill & Lilly 1991; Ledlow & Owen 1996; Zirbel 1997; Miller et al. 1999; Stocke et al. 1999; Wing & Blanton 2011; Gendre et al. 2013; Fujita et al. 2016; Shen et al. 2017), while very powerful FRIIs are more commonly found in rich environments at high redshift (Best 2000), although they have been found in clusters at low- z (e.g., Wing & Blanton 2011). Despite the distinct morphological divide between FRIs and FRIIs, recent analysis in Mingo et al. (2019) suggests that this difference in morphology does not follow the difference in the radio luminosity.

One unique type of radio AGN commonly found in clusters of all masses are bent, double-lobed radio sources (e.g., Blanton et al. 2000; Sakelliou & Merrifield 2000; Blanton et al. 2001; Wing & Blanton 2011, 2013; Paterno-Mahler et al. 2017; Silverstein et al. 2018; Garon et al. 2019; Golden-Marx et al. 2019; de Vos et al. 2021; Golden-Marx et al. 2021; Vardoulaki et al. 2021; Morris et al. 2022). The radio luminosity of bent radio sources typically straddles the FRI/FRII border (Blanton et al. 2000, 2001), and their unique bent morphology implies the presence of a dense gaseous medium (like the intracluster medium, ICM), which is necessary to create the ram pressure to bend the radio lobes (e.g., Owen & Rudnick 1976; O'Donoghue et al. 1993; Hardcastle et al. 2005; Morsony et al. 2013).

Low- z bent radio AGNs are preferentially found in rich cluster environments (e.g., Wing & Blanton 2011; Morris et al. 2022), and in galaxy clusters between 40% and 80% of the time depending on the richness, about a factor of 2 higher than straight radio AGNs (Wing & Blanton 2011). Outside of galaxy clusters and groups, low- z bent radio AGNs have been found in large-scale filaments between galaxy clusters (e.g., Edwards et al. 2010) and fossil groups, neither of which would necessarily be characterized by large overdensities of red sequence galaxies. The study of the environments of bent radio AGNs was extended out to high redshift using the Clusters Occupied by Bent Radio AGN (COBRA) survey (Blanton et al. 2015; Paterno-Mahler et al. 2017; Golden-Marx et al. 2019, 2021). Similar to their low- z counterparts, high- z bent radio AGNs have been found in rich, red sequence selected galaxy cluster candidates as well as poorer and underdense environments (Golden-Marx et al. 2019).

Bent radio AGN morphology is dependent on environment, which leads to the question of whether one can one differentiate the environment of bent radio AGNs at high- z based solely on radio characteristics? If so, does this allow us to better determine which bent radio AGNs to target for cluster searches, thus yielding a potential method of identifying new, lower-mass, high- z galaxy clusters and groups?

In this paper, we take advantage of newly available radio observations from the LOw-Frequency ARray (LOFAR) Two-metre Sky Survey (LoTSS) Data Release II (Shimwell et al. 2022) to attempt to answer these questions. We present two

samples of bent radio AGNs previously identified in the high- z COBRA survey, one in red sequence cluster environments and one in noncluster environments. We compare their radio properties, including the projected physical size, projected physical area, radio luminosity, and spectral indices.

The organization of the paper is as follows. We introduce the COBRA survey in Section 2.1 and our cluster and noncluster sample in Section 2.1.1. We discuss our archival radio observations and the measurements of the radio source properties in Sections 2.2 and 2.3. We compare the properties of bent radio AGNs in cluster candidates to those outside of clusters in Section 3. Lastly, we compare our findings to those of similar work probing the environment and properties of radio AGNs in Section 4. Throughout this work, we adopt a flat Λ cold dark matter cosmology, using $H_o = 70\text{ km s}^{-1}\text{ Mpc}^{-1}$, $\Omega_m = 0.3$, and $\Omega_\Lambda = 0.7$. Unless noted, all distances are given as projected distances, and all magnitudes are given as AB magnitudes. In this paper, we refer to high- z as the redshift range of the bent radio AGNs in this study, $0.35 < z < 2.35$, or similar redshifts (with the majority of the sources being at $0.5 < z < 1.5$). For the remainder of the paper, we refer to Golden-Marx et al. (2019) as GM19 and Golden-Marx et al. (2021) as GM21.

2. Data

In order to better characterize the properties of high- z bent radio AGNs as a function of environment, we need a high- z cluster sample and a high- z noncluster sample. For our analysis, we take advantage of bent radio AGNs identified as part of the high- z COBRA survey. To do this analysis, we use publicly available data from the Very Large Array (VLA) Faint Images of the Radio Sky at Twenty-centimeters (FIRST; Becker et al. 1995) and LoTSS DR2 (e.g., Shimwell et al. 2017, 2019, 2022). Our principal focus are the newly available LoTSS DR2 (Shimwell et al. 2022) data described in Section 2.2.

2.1. The High- z COBRA Survey

The high- z COBRA survey is a sample of 646 bent, double-lobed radio sources selected from the FIRST survey at $0.35 < z < 3.0$ (with the majority of AGNs having redshifts between $0.5 < z < 1.5$). These bent radio AGNs were originally identified as part of the sample of radio sources studied in Wing & Blanton (2011). Specifically, Wing & Blanton (2011) built on the visually selected sample of 384 bent radio AGNs from Blanton et al. (2000) to identify potential bent radio AGNs in the $\approx 10,000\text{ deg}^2$ FIRST survey. Unlike the visual-bent selected sample, Wing & Blanton (2011) used pattern recognition software from Proctor (2006) to identify 1546 three component radio sources in what they refer to as the auto-bent sample. Because FIRST was designed to cover the same portion of the sky as the Palomar Sky Survey, which is the same area of the sky observed by the Sloan Digital Sky Survey (SDSS), Wing & Blanton (2011) cross-matched the radio sources with SDSS DR7 to identify low- z host galaxies. As not every bent AGN had a host galaxy detected above the magnitude threshold ($m_r = 22.0$ mag) used in Wing & Blanton (2011), the 653 remaining sources became the initial high- z COBRA sample (e.g., Blanton et al. 2015; Paterno-Mahler et al. 2017). Of these 653 sources, 646 were successfully observed with the IRAC 3.6 μm band on Spitzer (PI Blanton).

Of these 646 sources, 41 are spectroscopically confirmed quasars from SDSS. Using this sample, Paterno-Mahler et al. (2017) cross-matched the $3.6\ \mu\text{m}$ images with radio observations from FIRST to identify potential host galaxies. Paterno-Mahler et al. (2017) then measured a single-band IR overdensity in both $1'$ and $2'$ regions centered on the radio AGN and found that $\approx 82\%$ of bent radio sources are in overdense environments, but only $\approx 29\%$ are found in rich cluster candidates, defined by a minimum of a 2σ overdensity relative to a background field.

To further determine if these bent radio AGNs are in cluster candidates, GM19 analyzed the population of red sequence galaxies surrounding each radio AGN in a subset of the COBRA sample. To estimate the color of the surrounding galaxies, GM19 introduced i -band observations of 90 of the fields hosting bent radio AGNs taken using the 4.3 m Lowell Discovery Telescope (LDT). These fields were chosen based on either a particularly clear bent radio morphology, the radio sources being identified as quasars in SDSS, or the overdensity measurements from Paterno-Mahler et al. (2017). Of these 90 fields, 38 were also observed in the r band on the LDT (see Table 1 in GM19 for the observed fields and exposure times).

All cluster candidate redshift estimates are based on the identification of the host galaxy. For the bent radio AGNs without spectroscopic redshifts, GM19 estimated a photometric redshift by comparing the colors of the host galaxies to spectral energy distribution (SED) models of quiescent galaxies across the redshift range $0 < z < 3.0$ generated from EzGal (Mancone & Gonzalez 2012), with an estimated redshift error of ± 0.1 (see GM19 for the input parameters of the models used; for the radio AGNs in this paper, the redshift, error, and source of that redshift are specified in Table 2). Of the 90 radio AGNs observed in the optical, GM19 estimated photometric redshifts or found an SDSS spectroscopic redshift for 77 AGNs. Although only four of the sources with spectroscopic redshifts are not quasars (and thus have a color typical of a red sequence galaxy), GM19 found strong agreement between their photometric redshift estimates using optical and IR photometry and sources with SDSS photometric or spectroscopic redshifts.

Using the available optical and IR observations, GM19 measured the color of the host galaxy identified in Paterno-Mahler et al. (2017) ($i - [3.6]$, $[3.6] - [4.5]$, and/or $r - i$) and compared it to the color of the surrounding galaxies. To identify potential red sequence galaxy cluster candidates, GM19 used a red sequence width of ± 0.15 from the color of the host in the $i - [3.6]$ or $r - i$ color, in agreement with the 3σ red sequence scatter found in spectroscopic studies at $z \approx 1.0$ (e.g., Blakeslee et al. 2003; Mei et al. 2006, 2009; Lemaux et al. 2012; Snyder et al. 2012; Cerulo et al. 2016). For the higher-redshift sources observed primarily with Spitzer photometry, GM19 used a slight modification to the well-studied Spitzer color cut (Papovich 2008; $[3.6] - [4.5] > -0.15$), which has been used to identify numerous high- z cluster and protocluster candidates (e.g., Wylezalek et al. 2013).

To estimate which fields host galaxy cluster candidates, GM19 used two distinct methods. First, they compared the number of red sequence galaxies within a $1'$ region centered on the radio AGN to the average number in a background field. Second, to account for all bent radio AGNs that are not at the centers of their cluster (e.g., Sakelliou & Merrifield 2000), GM19 measured the surface density of red sequence galaxies within the field of view (FOV) and chose the peak of that

distribution as the new cluster center. We use these red sequence surface density measurements and cluster centers in this paper.

Because cluster galaxies, especially at higher redshift, are not always red, GM19 also measured a combined overdensity, which accounts for the overdensity of red sequence galaxies, galaxies bluer than the red sequence, and galaxies redder than the red sequence. To scale this measurement, GM19 used observations with more statistically robust photometric redshifts from the Observations of Redshift Evolution in Large-Scale Environments Survey (ORELSE; e.g., Lubin et al. 2009; Hung et al. 2020), to measure the fraction of galaxies at a given redshift in the ORELSE survey that are on the red sequence or redder/bluer. For both overdensity measurements, the detection threshold is either an $m^* + 1$ magnitude limit (if $z < 1.0$), where the magnitude of an m^* galaxy is estimated via EzGal, or the magnitude limit of the Spitzer observations (21.4 mag). To prevent single galaxy detections from yielding a cluster candidate at the highest redshifts, GM19 require both a 2σ detection when centered on either the radio AGN or the estimated cluster center and that each cluster candidate have at least three red sequence galaxies within the $1'$ search region. From the sample with redshifts, GM19 identified 39 high- z bent radio AGNs in galaxy cluster candidates at $0.35 < z < 2.2$ (see GM19 for the red sequence and combined overdensities).

2.1.1. The High- z Cluster and Noncluster Sample

Although GM19 identified 39 red sequence galaxy cluster candidates, after further analysis of the bent radio AGNs, GM21 removed three radio AGNs from the sample due to uncertain morphology. As such, the initial high- z sample in this work consists of the 36 cluster candidates in GM21. To create a uniformly observed and analyzed comparison sample of high- z bent radio AGNs that are not in clusters, we return to the sample from GM19. Since certain characteristics depend on redshift (e.g., radio luminosity, projected physical size, projected physical area), we only select AGNs that either have a spectroscopically confirmed redshift from SDSS, a photometric redshift from SDSS, or an estimated redshift from GM19. As mentioned above, 39 bent radio AGNs with optical imaging and redshift estimates are found in cluster candidate environments, while 38 are not. Thus, using this sample, we select bent AGNs with no evidence of being in rich, red sequence cluster environments. These sources have the same available photometry as the cluster sample. Although GM19 does not identify any of these sources as cluster candidates based on the surrounding distribution of red sequence galaxies, we explore whether these radio AGNs are truly representative of a field sample in Section 4.1.

2.2. Radio Observations

While much of the pioneering work on large samples of bent radio AGNs in clusters was done using FIRST (e.g., Blanton et al. 2000, 2001; Wing & Blanton 2011, 2013), we are entering the LOFAR era of radio astronomy, and understanding the similarities in radio properties is vital to further characterizing bent radio AGNs. We chose LOFAR due to its sensitivity and because lower-frequency emission is more sensitive to older, less energetic synchrotron emitting electrons. This less energetic synchrotron emission may be more impacted by the cluster environment, making LOFAR the ideal tool for our

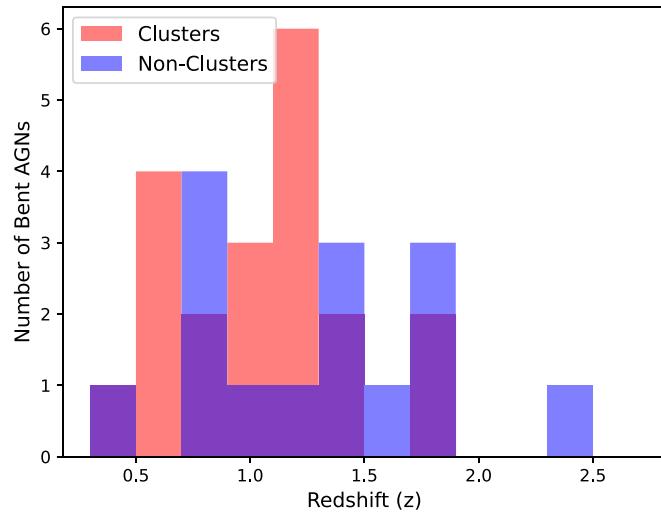


Figure 1. The redshift distribution of bent radio AGNs in the cluster (red) and noncluster (blue) sample. While we have more cluster sources at $0.5 < z < 1.5$ than noncluster sources, each sample spans a similar redshift range (with the exception of the highest-redshift noncluster source).

analysis. Specifically, we use the publicly available data from the LoTSS DR2 (Shimwell et al. 2022) observed in the 120–160 MHz range (with a median frequency of 144 MHz).

Of the 36 COBRA cluster candidates and 38 COBRA noncluster candidates, only four sources were observed with LoTSS DR1 (e.g., Shimwell et al. 2017, 2019). However, 37 COBRA sources from both samples were observed as part of LoTSS DR2. In total, 21 are in high- z cluster candidates and 16 are in the noncluster sample. Of the 21 high- z cluster candidates, we remove one source where we do not have a well-identified host galaxy (see Figures 2 and 3 for the LoTSS images and Figures 18 and 19 for Spitzer images with host galaxy identification). Additionally, of the 16 noncluster sources, we remove one radio AGN that is not actually a bent radio AGN, but two overlapping radio sources. For the remainder of this paper, the high- z cluster sample (noncluster sample) is these 20 (15) bent radio AGNs (see Table 2 for the full list of bent radio AGNs in each sample).

The goal of this paper is to study the impact of environment on bent AGNs, not redshift evolution. While a true analysis of redshift evolution is beyond the scope of this paper due to the small sample, we show the redshift distribution of each sample in Figure 1. As shown in Figure 1, despite there being more bent radio AGNs in the cluster sample at $0.5 < z < 1.5$ than in the noncluster sample, both samples include radio AGNs that cover a similar redshift range. However, there is likely some redshift evolution influencing the observed properties, especially the environment. Specifically, the density of radio AGNs in clusters increases with redshift (e.g., Galametz et al. 2009; Mo et al. 2018, 2020) and the cluster environment has greater impact on the opening angle of bent AGNs due to the ICM density decreasing as a function of increasing redshift (Vardoulaki et al. 2021). Additionally, the luminosity functions of low- and high-luminosity radio AGNs peak at different redshifts (Hardcastle & Croston 2020). Given the similar redshift distributions, we do not expect redshift evolution to be dramatically different between the two samples.

Of note, the highest-redshift source in the noncluster sample is at $z = 2.346$, while the highest-redshift source in the cluster sample is at $z = 1.818$ (both are SDSS spectroscopically

confirmed quasars). This difference in redshift is not necessarily an evolutionary factor, where bent AGNs are only found in clusters at $z < 2.0$. GM19 identified one cluster candidate at $z = 2.18$, which to date has not been observed by LoTSS. Additionally, we find a slightly higher fraction of the bent radio AGNs in the noncluster sample are SDSS-identified quasars (6 of 15 [40.0 %] versus 4 of 20 [25.0%]; see Table 2). Because quasars are typically more energetic than other radio AGNs and have been shown to have slightly different properties in terms of their locations within high- z clusters (e.g., Mo et al. 2018), this difference may impact our analysis.

2.2.1. Radio Source Projected Physical Size

We aim to determine if the size and asymmetry of bent radio AGNs is impacted by the cluster environment. We estimate the projected physical size of each bent radio AGN in a similar manner to Moravec et al. (2019, 2020), who estimated the largest physical extent by identifying the two most distant points along 4σ contours of their radio observations. Since the “C” shape of bent radio AGNs means that the two lobes might be closer to one another than to the core, we treat each lobe separately and measure the largest projected physical extent from the radio source core (identified from FIRST/Spitzer in Paterno-Mahler et al. 2017 and GM19) to each lobe and combine these two distances to get our measurement of the projected physical size (see Figures 2 and 3 for examples showing the rays tracing the projected physical size, and see Table 2 for our measurement of the projected physical size and the angular sizes of each lobe). For bent radio AGNs with less complicated morphology, this is approximately the total extent of the radio source if it were straight. However, for the more complex sources, this measurement tends to underestimate the true physical extent.

We determine our detection criteria by measuring the rms noise in each LoTSS image. Although we only show $1/3 \times 1/3$ cutouts in Figure 2 and 3, we did our analysis using $5' \times 5'$ cutouts centered on the radio AGN. We estimate the rms by analyzing a 50 pixel \times 50 pixel region in each LoTSS cutout ($\approx 1/25 \times 1/25$). This region is near the radio source and devoid of other bright radio sources. We report the rms values of each field in Table 3. Based on the appearance of radio artifacts, as well as the high level of detection in the LoTSS data, we use a 10σ threshold for the minimum detection for each bent radio AGN. This value is based on the rms noise of each image to account for slight differences in the local background.

As size measurements are highly dependent on the detection threshold, we estimate the error in our projected physical sizes by measuring the difference in the size of each radio lobe if we instead use 9σ or 11σ as our minimum threshold. To more thoroughly account for this error, we do not extend the ray between the location of the maximum projected distance at 10σ to the 9σ and 11σ contours. Rather, we remeasure the location of the greatest physical extent at 9σ and 11σ and measure the projected physical size at those contour values. We then add this measurement error in quadrature with our estimate of our observational error, which we assume is half of the beam radius ($3''$ for LoTSS) to account for the total error. To convert our projected angular size into a projected physical size, we multiply by the angular diameter distance, which we calculate using astropy (Astropy Collaboration et al. 2013, 2018). We estimate the error in the angular diameter distance following

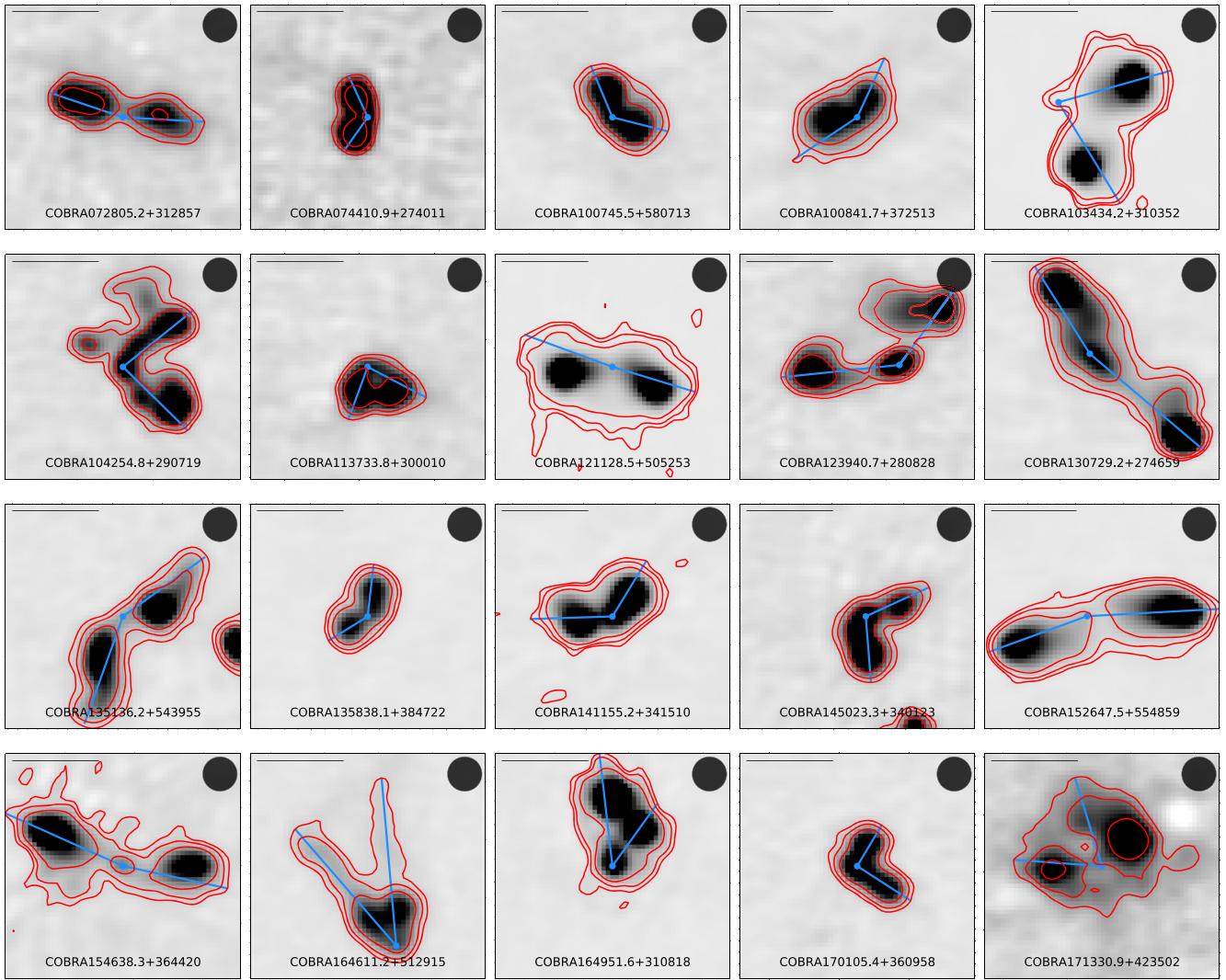


Figure 2. Cutouts of the 20 bent, double-lobe radio sources in the COBRA cluster candidate sample observed by LoTSS DR2. While most are $1/3 \times 1/3$, some are slightly larger to encompass the extent of the entire radio source. The three red contours show 10σ , 20σ , and 50σ contours based on the measurement of the rms noise in each image. The blue lines show our estimate of the projected physical size of each radio lobe, and the filled blue circle shows the location of the AGN’s core identified from both the optical/IR host and the FIRST images (see Figure 18 for the optical/IR host galaxies). The black circle shows the beam size, and the black line shows $0/5$.

GM21 and measure the difference in the angular diameter distance based on the error in the redshift (see Table 2). We then combine these values with the measurement error to get the total error in the projected physical size. Because our combined error is larger than the $0.^{\circ}2$ astrometric uncertainty reported in Shimwell et al. (2022), we do not account for this additional source of error.

To verify the validity of these measurements, we examine each source by eye. The extent of most sources is well estimated by our measurement of the projected physical size. However, a few sources have greatest physical sizes that are impacted by small features in the 10σ contour creating rays that do not trace the structure as determined by eye (e.g., COBRA104254.8+290719, COBRA154638+364420). For example, COBRA104254.8+290719 has an additional radio feature to the left of the upper lobe, which was initially identified as the location of the greatest radio extent. While these features are real (and the feature for COBRA104254.8+290719 appears, though much fainter, in the FIRST image as well), we ignore these features when measuring the projected

physical size and use the greatest extent from the main lobe as shown in Figure 2. Additionally, for COBRA104541+282028 (see Figure 3), the radio contours trace two separate lobes with no identified core. For this source, we estimate the size as the position from the core, which is the location of the host galaxy identified in Paterno-Mahler et al. (2017), to the farthest extent of the radio lobes. Thus, the physical extent extends beyond the actual size traced by the radio lobes, which may add an additional source of error.

Although we treat the projected physical size as an intrinsic property of the radio source, it is highly dependent on the redshift and our measurement of corresponding noise levels. While we use a 10σ clipping for each radio AGN, the surface brightness of the radio source is dependent on the redshift, and thus, the level of that 10σ clip relative to the radio lobes is also impacted. Since the majority of our sources have lobes with brighter edges, they are not as impacted as sources such as COBRA164611.2+512915 and COBRA171330.9+423502 in Figure 2 and COBRA090745.5+382740 and COBRA132903.2+253110 in Figure 3, which have clearly visible faint tails. To determine the

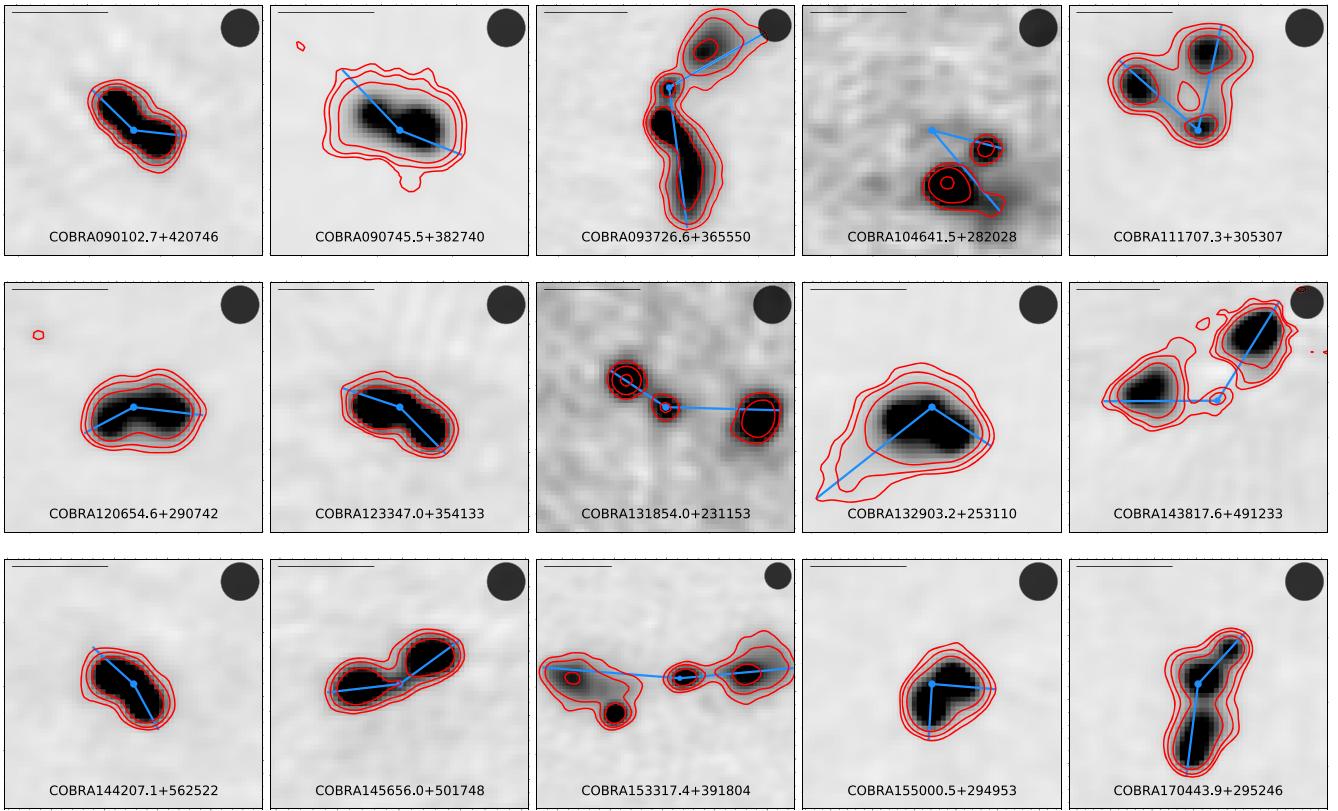


Figure 3. Cutouts of the 15 bent, double-lobed radio sources in the noncluster sample observed by LoTSS DR2. While most of the cutouts are $1'3 \times 1'3$, some of the images are slightly larger due to the larger angular size of the radio AGNs. All contours, lines, and circles are the same as in Figure 2 (the location of the core for COBRA104641.5+282028 comes from the optical/IR host and the FIRST contours; see Figure 19 for the host galaxies).

effect on the system as a whole, we examine the total angular size as a function of redshift using a Spearman test. For the cluster sample, we find a weak correlation with redshift and no evidence to reject the null hypothesis ($r_s = -0.25, p = 0.29$). For the noncluster sample, we find a moderate trend, with no evidence to reject the null hypothesis ($r_s = -0.42, p = 0.12$). Given the smaller sample size and that the highest-redshift source (one of the smallest sources) is in the noncluster sample, we remove this source to determine if it is an outlier. In this case, we find a weak trend with no evidence to reject the null hypothesis ($r_s = -0.310, p = 0.281$). While our sample is not statistically impacted by the redshift dependence of the measurement, it may slightly bias some of our results and should be taken as another potential source of error when we discuss the impacts of physical size and the asymmetry of radio lobes as a function of environment.

2.2.2. Radio Source Projected Physical Area

Although the projected physical size is a common measurement of radio AGNs (e.g., Moravec et al. 2019, 2020; Shen et al. 2020), the physical size is only a one-dimensional measurement. The cluster environment might limit the degree of collimation within each radio lobe, which would not necessarily be accounted for in a length measurement. As such, we estimate the projected physical area of each radio lobe to account for the total volume of the lobes. We measure the projected physical area by measuring the area enclosed within each 10σ contour using Green’s Theorem, which transforms the line integral over the curve of our contour into a double integral over the region within the contour. We convert this to

the projected physical area by multiplying the area within the contour by the angular diameter distance squared (see Table 2 for the total projected physical area measurements).

As we are interested in probing asymmetry between the bent lobes of each AGN, we estimate a fiducial divide centered on the radio core and extending to the 10σ contour to separate the radio lobes (see Figure 4). While this allows us to measure lobes separately, we do not treat the central core region as a separate component (this region becomes evident in some sources when we look at the spectral index—see Section 2.3), but instead include it equally as part of the area of each lobe. Since the core region is relatively symmetric, we believe this is the best solution in estimating lobe asymmetries.

Unlike the projected physical size, for the projected physical area, we only measure the extent within the 10σ contours of the main lobes (for some of the radio sources, such as COBRA104254.8+290719, there are additional associated components, which we do not include). As discussed for our measurement of the projected physical size, there is one example where no emission is detected at the core at the 10σ level and multiple sources that are detected as separate components (See Figure 2). While our measurement of the largest projected physical size accounts for the distance regardless of whether or not the entire source is contained within the 10σ contour (including the core), our measurement of projected physical area only includes the area of the components, not any space connecting the core to the lobes. Thus, our measurements of the projected area may be a slight underestimate, as there should be a radio jet connecting the regions.

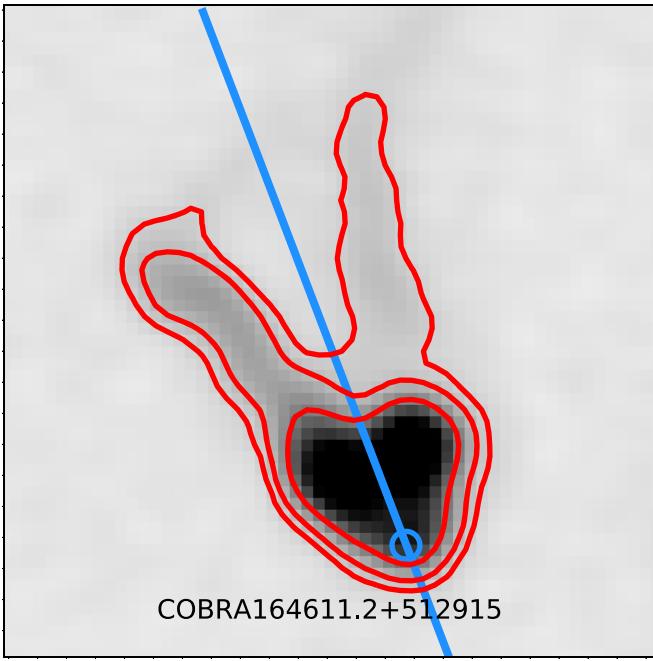


Figure 4. Radio contours for COBRA164611.2+512915. Here we show our fiducial divide of the radio source (the blue line) through the radio core that we use to divide the radio AGNs into lobe 1 and lobe 2. As in Figure 2, the blue circle shows the AGN core, and the red contours show 10σ , 20σ , and 50σ values.

We estimate the error in this measurement following the same approach as in Section 2.2.1. Using 9σ and 11σ contours, we measure the area of each lobe and compare it to the 10σ value, which we similarly combine in quadrature with half of the area of the radio beam. We then combine this error with the error measurement for the angular diameter distance to account for the total error.

Similar to the projected physical size, the projected physical area is subject to the redshift dependence of surface brightness relative to our 10σ flux density threshold and is not an intrinsic property of the radio source. Again, we do a Spearman test to compare the angular area as a function of redshift and find a very weak correlation and no evidence to reject the null hypothesis within the cluster sample ($r_s = -0.18$, $p = 0.43$). We again find a moderate correlation among the sources in the noncluster sample, with no evidence to reject the null hypothesis ($r_s = -0.44$, $p = 0.10$). Like the projected physical size, the slight redshift dependence in the noncluster sample may be due to the highest-redshift source being one of the smallest sources. When we do the Spearman test without this source, we find a weaker trend with no evidence to reject the null hypothesis ($r_s = -0.36$, $p = 0.20$). While we see no statistical trends with redshift, it is possible that our measurements of the projected physical area and the asymmetry of the radio lobes are impacted by this redshift effect. As such, this should be treated as an additional source of error in our analysis.

2.2.3. Radio Source Opening Angle

Although our measurements of the projected physical size of each radio lobe yield a first-order estimate of the opening angle of each bent radio AGN, there are a number of sources, including COBRA171330.9+423502, where the projected

physical extents do not trace the peak regions of flux and thus yield an inaccurate representation of the opening angle. To better capture the opening angle for all bent radio AGNs, we measure the angle between the radio source core (as identified in GM21) and the brightest region in each lobe (see Figure 5).

We quantify this opening angle, following the convention in GM21:

$$\text{Opening Angle} = \arccos\left(\frac{(H)^2 - (L_1)^2 - (L_2)^2}{-2(L_1)(L_2)}\right), \quad (1)$$

where H is the ray connecting the two brightest pixels in each lobe, and L_1 and L_2 are the rays between the core and the brightest pixel in each lobe, respectively. Although we define lobe 1 and lobe 2 and treat them separately throughout the majority of this paper, they are arbitrarily defined, meaning any trends between properties of lobe 1 and lobe 2 simply show a degree of asymmetry, not necessarily a defining characteristic of each lobe. Similar to GM21, we estimate the error in the opening angle using the error in the relative position of each component.

As shown in Figure 6, by comparing the opening angle measurements from GM21 from FIRST with the opening angle measurements from LoTSS (see Table 2 for the full list of opening angles), we find that there is relative agreement between the values, with all but four sources having a measurement difference of less than 10%. We individually examine these four sources with the largest difference in the opening angle. For COBRA164951.6+310818, the major offset is due to an error misidentifying the radio source core (and thus the opening angle) in GM21. Similarly, for COBRA145023.3+340123, the difference comes from differences in the identified lobes. As shown in Figure 7, in the FIRST image, the component to the west of the core, shown as a separate component in FIRST, is much fainter and was not included in the FIRST catalog. Thus, we instead identified the second lobe as what appears to be a secondary lobe feature in the eastern lobe in the LoTSS image. Although we note that the western component in COBRA145023.3+340123 does appear to be a radio lobe, the lack of inclusion in the FIRST catalog is intriguing. While other radio AGNs like COBRA164611.2+512915 have additional radio emission extending beyond the brightest components of the main lobes, no other source has a previously unreported feature. As such, we caution that this new lobe may be a co-spatial source and requires further analysis to accurately trace the opening angle of this radio source. For the third source, COBRA113733.8+300010, which is in the only spectroscopically confirmed cluster in our sample (Blanton et al. 2003), the difference is due to the highly curved nature of the lobes, which create differences based on the location of the peak brightness in the lobes. For the fourth source, COBRA170105.4+360958, the offset is slight ($\approx 12\%$), and again likely due to small differences in the location of the peak brightness.

2.2.4. Radio Luminosity

Beyond the visual morphology, we are also interested in probing the differences in the total radio luminosity of each radio source to see if any differences exist among the populations inside and outside of clusters. To measure the

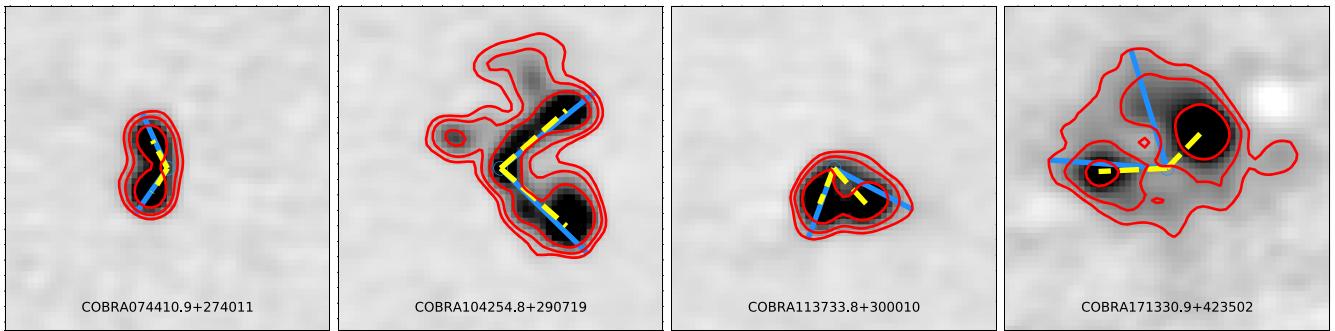


Figure 5. Examples of our measurement of the opening angle for the LoTSS images. The red contours are the same contours as in Figure 2. The yellow dashed lines show L_1 and L_2 from Equation 1, while the blue lines show our estimate of the opening angle using the largest physical extent of each radio source. In some cases, the opening angles are nearly identical, while most slightly differ. However, some give a dramatically different value, as is the case for COBRA171330.9+423502, in the far right panel.

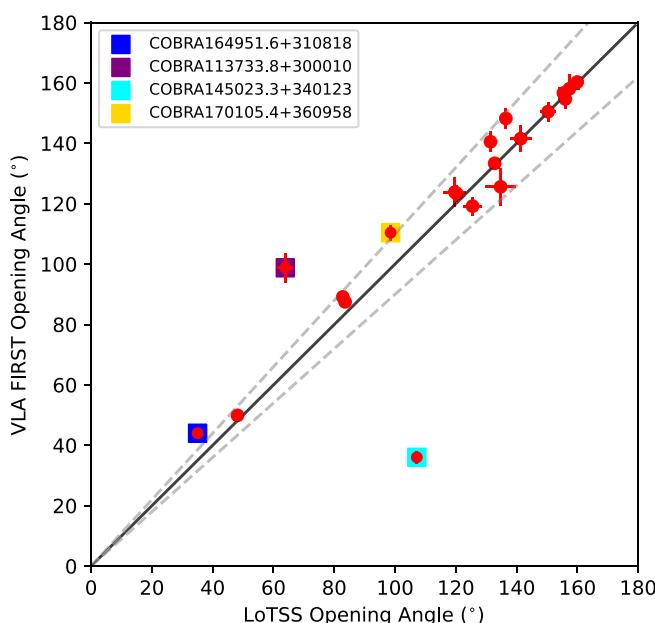


Figure 6. A comparison of the opening angle measurements from FIRST (GM21) with the opening angle measurements from LoTSS DR2 for the cluster sample. The black lines shows a one-to-one trend, while the two gray dashed lines highlight a 10% difference between the opening angles. We overlay colored squares on all bent AGNs where the opening angles differ by more than 10%. For the majority of our sources, these measurements are very similar.

radio luminosity, we follow Moravec et al. (2019, 2020) and GM21 and use Equation (2),

$$L_{144 \text{ MHz}} = 4\pi D_L^2 S_{144 \text{ MHz}} (1 + z)^{\alpha - 1}, \quad (2)$$

where $L_{144 \text{ MHz}}$ is the radio luminosity (W Hz^{-1}) at 144 MHz, the median frequency of the LoTSS survey, D_L is the luminosity distance (Mpc), calculated with Astropy using the redshift estimates from GM19, α is the spectral index calculated by measuring the integrated spectral index of the entire source (see Section 2.3), and $S_{144 \text{ MHz}}$ is the integrated radio flux density (mJy) of each radio source. Although Shimwell et al. (2022) reported flux densities for these radio sources, to more accurately capture the flux density of the radio sources for our analysis, we use the CASA tool `imstat` to measure the flux density across the radio source (above the 10 σ threshold, as shown in the radio contours in Figures 2 and 3).

For our bent radio AGNs in cluster candidates, we find that the radio luminosities range from $1.7 \times 10^{25} \text{ WHz}^{-1}$ to $4.1 \times 10^{27} \text{ WHz}^{-1}$ (see Table 3 for the full list of radio luminosities). Compared to the measurements to the radio luminosities from GM21, we find that the three most powerful radio sources in the cluster sample are three of the four most powerful radio sources reported in that paper (the fourth-brightest bent AGN has not been observed as part of LoTSS). Similarly, we find that for the bent radio AGNs in the noncluster sample, the radio luminosities range from $0.8 \times 10^{25} \text{ WHz}^{-1}$ to $3.3 \times 10^{27} \text{ WHz}^{-1}$. Based purely on the radio luminosity, we do not see any major differences between the two samples.

2.3. Spectral Index

To more completely probe the energetics of each bent radio AGN, we measure the integrated spectral index of the radio emission across the radio source. Because the SED of an AGN is dominated by synchrotron emission, it can be estimated by a smooth power law of the form ν^α . The value of α , the spectral index, is typically of the order -0.7 to -0.8 for extended radio sources (e.g., Kellermann & Owen 1988; Sarazin 1988; Condon 1992; Peterson 1997; Lin & Mohr 2007; Miley & De Breuck 2008) and can be calculated by measuring the flux density in different wave bands.

Shimwell et al. (2022) measured a spectral index for sources identified in the LoTSS DR2 catalog by dividing the wide LOFAR band into two bands, and noted that multiple factors led to a large degree of uncertainty, with error in the value of the order $\gtrsim 0.2$. To better probe the energetics of the core and lobes, we measure the spectral index by combining the LoTSS data with publicly available data from FIRST (Becker et al. 1995; the same data used in GM19 and GM21). Using these two bands, we cover a frequency range of 144 MHz (LoTSS) to 1.44 GHz (FIRST). To accurately compare each bent radio AGN in each wave band, we use the CASA tool `imsmooth` (McMullin et al. 2007) to smooth the FIRST images to the angular resolution of LoTSS (6''); the angular resolution of FIRST is 5''. We then use the CASA tool `imregrid` to rescale the FIRST images to the pixel scale of LoTSS (1.5'' per pixel; the pixel scale of FIRST is 1.8'' per pixel). Using these recalibrated FIRST images, we measure the rms noise for the FIRST images using a 50 pixel \times 50 pixel region as done for our previous LoTSS analysis.

While our radio AGNs are very bright relative to the background in the LoTSS images, the same radio AGNs are not as bright relative to the background in the FIRST images (see

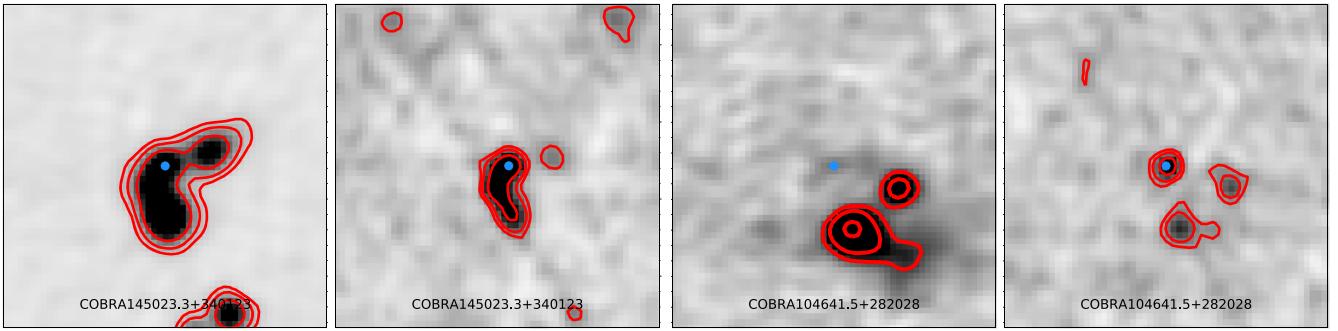


Figure 7. Examples of the two bands used to create spectral index maps. The left two images show LoTSS and FIRST images of COBRA145023.3+340123, while the right two images are from COBRA104641.5+282028. In the LoTSS images, the red contours show the 10σ , 20σ , and 50σ contours, while in the FIRST images, the red contours show 3σ , 5σ , and 10σ . The blue circle identifies the location of the core. Each FIRST image is smoothed and regressed to match LoTSS (6" resolution and $1''/5$ pixels). We specifically highlight two sources with slight differences in the components. For COBRA145023.3+340123, the west component is not as strongly identified in FIRST, while for COBRA104641+282028, the core is not detected at the 10σ level in LoTSS.

Figure 7). To determine the spectral index, we measure the ratio of the flux densities in the LoTSS and FIRST images. To best account for the flux density in the same spatial region of the radio source, we measure the flux density within the 3σ contours based on the FIRST images for both the FIRST and LoTSS observations. As shown in Figure 7, the 3σ FIRST contours generally fall within the 10σ LoTSS contours. However, for a small number of sources, the 3σ FIRST contours extend outside of the LoTSS contours. For those sources, we measure the flux within the 3σ FIRST contours in each band to best account for the flux of the radio source. As these regions are generally only slightly offset from the LoTSS 10σ contours, the FIRST contours should accurately probe the spectral index. The exception is the core in COBRA104641.5+282028, which is detected at the 3σ level in FIRST and not at the 10σ level LoTSS (this field is flagged in Table 3).

To calculate the spectral index, we first write the flux density, F , as:

$$F = A\nu^\alpha, \quad (3)$$

where A is the amplitude, ν is the rest-frame central frequency of the band, and α is the spectral index. We calculate the spectral index by taking the ratio of the two flux densities, noting that the amplitude is the same for the same radio source. Thus, the spectral index has the form:

$$\alpha = \frac{\log(\frac{F_1}{F_2})}{\log(\frac{\nu_1}{\nu_2})}, \quad (4)$$

where F_1 is the flux density (mJy) in LoTSS, F_2 is the flux density (mJy) in FIRST, ν_1 is the frequency of LoTSS, and ν_2 is the frequency of FIRST.

We measure the error in the spectral index from the error in the radio flux density at each frequency following the methodology to Di Gennaro et al. (2018a). For the LoTSS and FIRST images, we follow Shimwell et al. (2022), and assume a 10% error in these flux density measurements. We add this error in quadrature with the error due to the rms noise (see Table 3) to estimate the total error in the flux density. We then sum these errors in quadrature following Equation (1) in Di Gennaro et al. (2018a) to get the error in the spectral index.

3. A Comparison of the Physical Properties of Bent AGNs inside and outside of Cluster Candidates

3.1. Comparing Size and Asymmetry in Radio Lobes

To understand the impact of environment on bent radio AGN morphology beyond the opening angle (e.g., Hardcastle et al. 2005; Morsony et al. 2013; Garon et al. 2019; Vardoulaki et al. 2021), we compare the distribution of radio source sizes and areas within each sample. As shown in Figure 8 and Table 2, bent radio AGNs in the cluster sample span a similar range of projected physical sizes as their noncluster counterparts (256–724 kpc for the cluster sample and 275–838 kpc in the noncluster sample). Similarly, we find relative agreement between the projected physical areas of bent radio AGNs in the cluster sample and their noncluster counterparts (0.029 – 0.229 Mpc 2 for the cluster sample and 0.027 – 0.157 Mpc 2 for the noncluster sample).

To probe differences in environment, we compare the projected physical size and projected physical area of each lobe to one another (see Figure 9). We find that the majority of bent radio AGNs appear relatively symmetric (>80%), regardless of their large-scale environment. To avoid bias from size differences in larger radio AGNs, the right-hand panels of Figure 9 show histograms of the ratios of the projected physical size and projected physical area of each radio lobe. Although we treat lobe 1 and lobe 2 separately, these ratios show the smallest lobe over the largest lobe, regardless of which is lobe 1. Each right-hand plot confirms that the majority of sources have lobes with similar sizes and areas. Although both samples show sources with a similar degree of symmetry in terms of the projected physical size and projected physical area, as highlighted in Figure 9, the two sources where the ratios are the lowest are in the noncluster sample. Given the similarities, it is unclear if the environment or radio luminosity play any role in the asymmetry or size of bent AGNs (see Section 4.1), or if instead, the environments of these two samples are more similar than the analysis in GM19 suggests (see Section 4).

As discussed in GM21, all measurements in regards to the size, area, and opening angles of bent radio AGNs are subject to projection effects. Since it is difficult to quantify all possible viewing orientations of bent radio AGNs, we note that differences in both the projected physical size and the projected physical area might be a measure of the impact of viewing angle on these measurements rather than differences in the environment. Thus, sources with slight differences in the

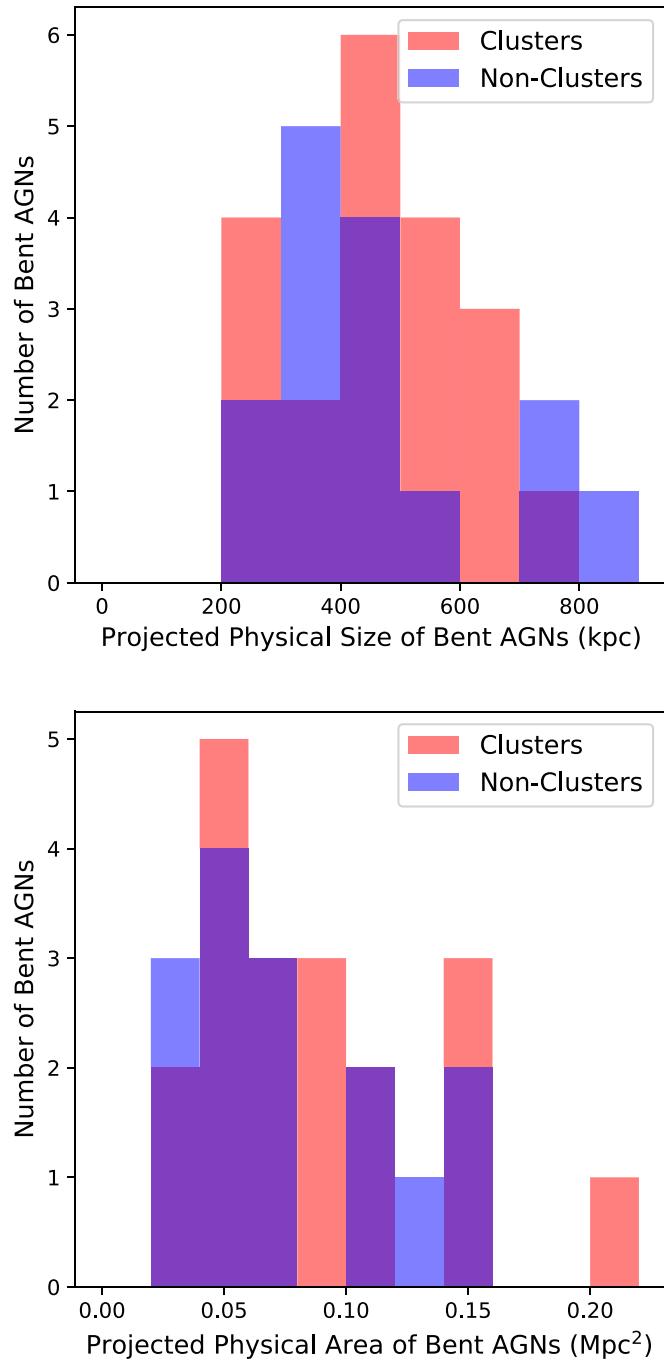


Figure 8. The projected physical size and area distributions of the bent radio AGNs in the cluster and noncluster sample. We see that each sample spans a similar range of sizes and areas.

projected physical size and projected physical area of the radio lobes may be indicative of similarly sized radio lobes viewed at differing orientations rather than slightly different environments or the effects of cluster weather and buoyancy.

For the remainder of the paper, our analysis of the size of bent radio AGNs will focus on the projected physical area. Although both measurements show a similar degree of symmetry, there are a number of sources (see Figures 2 and 3) where the rays estimating the projected physical size do not encapsulate the overall shape of the radio AGN and would require a more complex measurement to account for the unique morphology. As such, a more complex measurement of the

projected physical size, one that accounts for the radio flux density, frequency, sensitivity, and resolution to better trace the complete shape of the radio source and the redshift to trace the flux density cutoff, is needed to accurately map the length of bent radio AGNs. While imperfect, our measurement of the projected physical area better accounts for these factors.

3.2. Differences in Radio Luminosity

As discussed in Section 2.2.4, we measure the radio luminosity of each bent radio AGN using the flux density within our 10σ LoTSS contours. To look for differences between the radio luminosity of these two populations, we plot the radio luminosity ($L_{144 \text{ MHz}}$) as a function of the projected physical area of each radio source. As shown in Figure 10, we see a similar relationship between these populations. Statistically, the cluster sample (shown in red) shows a moderate trend with very strong evidence to reject the null hypothesis ($r_s = 0.630$, $p = 0.003$), with the most powerful radio AGNs being the largest. However, the noncluster sample (shown in blue) shows a weak to moderate trend with no evidence to reject the null hypothesis ($r_s = 0.371$, $p = 0.173$). The difference in the strength and statistical robustness of the correlations may imply that the cluster environment plays a minor role in shaping the trend between radio AGN size and luminosity, likely with respect to the overall size of the radio AGN. Alternatively, it may be indicative of the smaller noncluster sample size and the lack of the most luminous radio AGNs in that sample ($> 10^{27} \text{ W Hz}^{-1}$) weakening the overall trend (see Section 4.2 for further discussion).

To examine whether the difference in the strength of the statistical trends between the projected physical area and radio luminosity imply differences between the populations of bent AGNs, we run a Kolmogorov–Smirnov (K-S) test to compare the two quantities. In both cases, we find strong evidence that these samples are drawn from the same population ($p = 0.834$ for the radio luminosity and $p = 0.750$ for the projected physical areas). As the only underlying difference between these samples is the measurement of the environment from GM19, this is further evidence that either the role of the environment is marginal and only affects the correlation between the two values or that even though GM19 measured differences in the overdensity of red sequence galaxies, these sources are not in entirely different environments. Specifically, some of these noncluster systems, which have positive overdensities (but below the 2σ threshold), could be similar to the population of bent radio AGNs that are offset from clusters but still in local overdensities (Garon et al. 2019), and thus may be examples of poorer galaxy groups (e.g., Morsony et al. 2013; Vardoulaki et al. 2021).

3.3. Verification of the Opening Angle–Richness Correlation in Clusters

Using the larger COBRA high- z sample, GM21 found that richer galaxy clusters tend to host narrower bent AGNs. Assuming similar jet properties and the lack of projection effects, this would require a denser ICM for a greater degree of bending (e.g., Hardcastle et al. 2005; Morsony et al. 2013). We compare the opening angles measured with LoTSS to the combined overdensity defined in GM19 to determine the impact of the frequency and angular resolution on the observations on this correlation (see Figure 11).

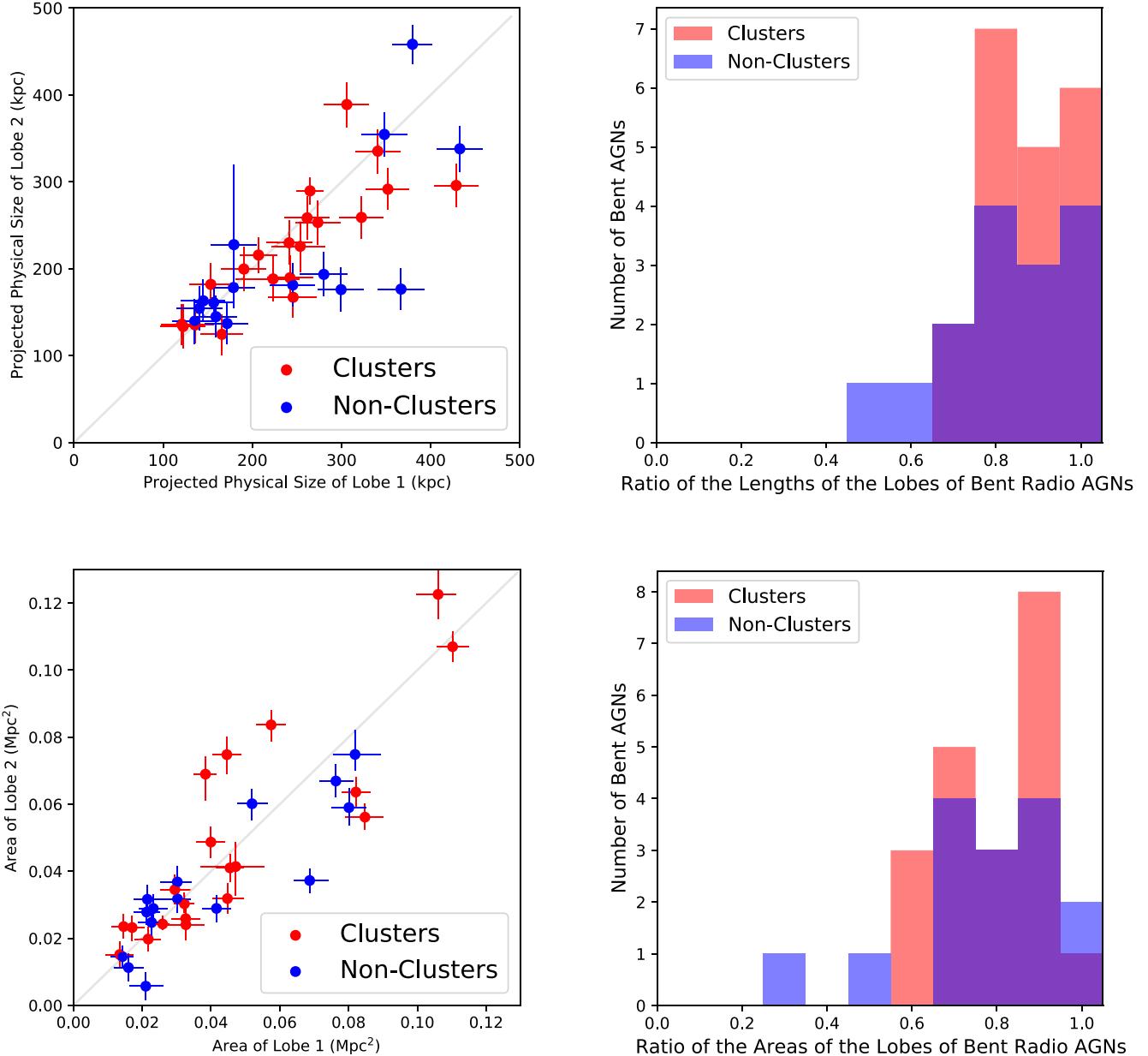


Figure 9. Probes of the symmetry of the lobes of bent radio AGNs. The top panels show the projected physical size (kpc) of each lobe when compared to the opposite lobe. The left-hand side is a direct comparison of the projected physical size, while the right-hand side shows a histogram of the ratios of the lengths of each lobe. For the ratio, we compare the smallest lobe length and largest lobe length regardless of which lobe is lobe 1 or lobe 2. Red points are bent AGNs in the cluster candidate sample. Blue points are bent AGNs in the noncluster sample. A gray dashed line in each of the left-hand panels shows a one-to-one trend. The bottom panels similarly show the projected physical area (Mpc^2) of each radio lobe when compared to its opposite lobe. It follows the same conventions as the top panel with the left side showing a direct comparison, and the right side showing a histogram of the ratio of the areas of each lobe. Both plots comparing the radio lobes indicate that most of the lobes have similar projected physical sizes and areas when compared to their opposite lobe. Although no overall trend with the total projected physical size or area is found, when we compare the ratio of the projected physical sizes and projected physical areas, the two most asymmetric sources are in noncluster environments.

Unlike the statistically robust trends identified in GM21, in our analysis in Figure 11, the correlation is not as statistically robust. Using a Spearman test, our trend for the $m^* + 1 - [3.6]$ sample shows a moderate trend, but little to no evidence to reject the null hypothesis ($r_s = -0.571, p = 0.180$). As highlighted in Section 2.2, although the LoTSS observations are excellent at tracing the overall larger scale of the radio emission of a bent radio AGN, the lower frequency means that LoTSS data is less effective at tracing the most energetic regions of radio emission. As the opening angle is estimated by measuring the curvature based on the brightest regions of the lobes and the

core, LoTSS data may introduce noise, which would impact our results.

To better estimate the biases in our measurements, we recompare the opening angles measured in FIRST from GM21 (we correct for the misidentified opening angle in GM21, which appears as an outlier to the trend) to the same combined overdensities (see the bottom panel of Figure 11). Unlike the larger sample in GM21, we find similar results in terms of both the strength and statistical robustness of the trend based on the Spearman test. However, for the FIRST measurements, we find that for the entire cluster candidate sample, and none of the

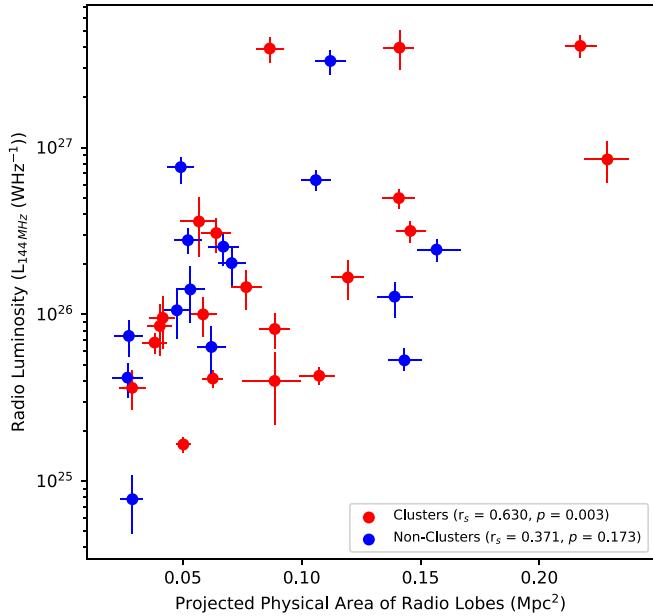


Figure 10. The radio luminosity of bent radio AGNs ($L_{144 \text{ MHz}}$) as a function of the total projected physical area of each bent radio AGN. The cluster candidates are shown in red, and the noncluster candidates are shown in blue. We find a moderate trend with strong evidence to reject the null hypothesis for the cluster sample, but only a moderate trend with no evidence to reject the null hypothesis for the noncluster sample.

subsamples, we see a moderate correlation with weak evidence to reject the null hypothesis ($r_s = -0.400, p = 0.082$). As these values imply only a slightly stronger and more statistically robust correlation than the LoTSS values, it would appear that the smaller sample size in this analysis may have a more dramatic effect than sensitivity, angular resolution, and wavelength coverage.

Although we do not find a statistical correlation, we considered the noncluster sample and looked at the median value of the opening angle and combined overdensity (for those with a combined overdensity measurement) and found that the median values were located at the upper-left portion of Figure 11. This adds further evidence that the opening angle is caused in part due to the environment (e.g., Hardcastle et al. 2005; Morsanyi et al. 2013; Garon et al. 2019).

3.4. Characterizing the Spectral Index of Bent Radio AGNs

Because of the angular resolution of LoTSS relative to the overall size of the bent sources, we only analyze the spectral index of the entire radio source, similar to Mahony et al. (2016), de Gasperin et al. (2018), and Dabhadé et al. (2020), or the core/lobe regions, similar to Di Gennaro et al. (2018b; see Figure 12 and Table 3). While we previously measured the asymmetry of radio lobes in Section 3.1, the coverage of the spectral index is not uniform across the two lobes for all of our sources, so we look at the lobes as a combined unit relative to the core as to not bias our results. Although a detailed accounting of the spectral index across the radio source is beyond the scope of this paper, we do present spectral index maps and error maps in Appendix B (see Figures 20, 21, 22, and 23).

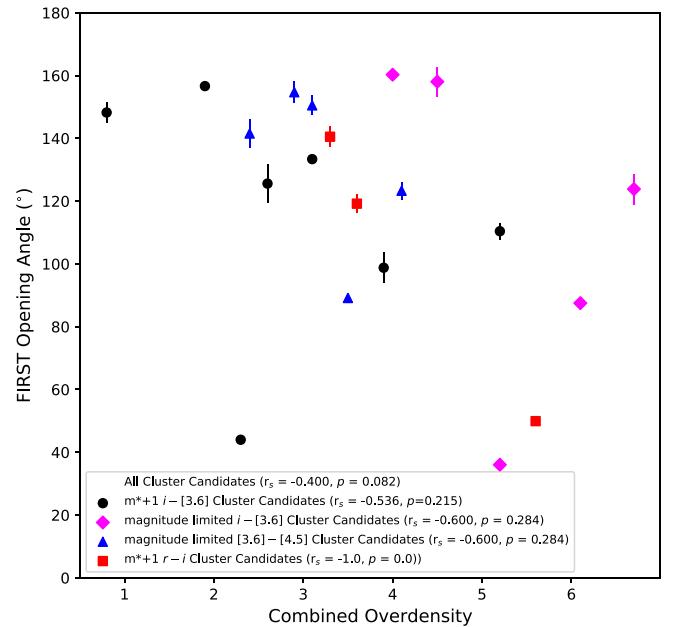
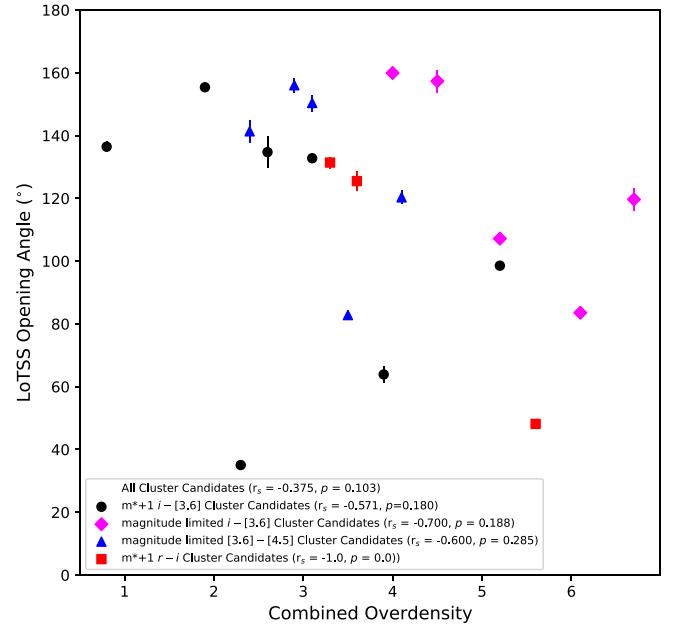


Figure 11. The LoTSS opening angle as a function of the combined overdensity (top panel) and the FIRST opening angle as a function of the combined overdensity (bottom panel). Unlike the trend shown in GM21, we find slightly weaker trends between the samples using the LoTSS images and very weak to no statistical evidence to reject the null hypothesis. However, we do find a moderate trend with weak evidence to reject the null hypothesis in the FIRST sample.

3.4.1. Total Spectral Index

As shown in the top panel of Figure 12, the total spectral indices of these sources span a similar range of values regardless of the environment. However, we find a slight difference in the peak of the distribution of values, with the median of the cluster sample being -0.76 ± 0.01 and the median of the noncluster sample being -0.81 ± 0.02 . Based on our error in the median, we find an $\approx 1.5\sigma$ offset between the median values of the total spectral index (see Table 1). This difference is further emphasized in Figure 12, where we find 10

Table 1
Median Values of the Mean Spectral Index

Sample	Total Spectral Index	Core Spectral Index	Lobe Spectral Index
Cluster Sample	-0.76 ± 0.01	-0.51 ± 0.03	-0.83 ± 0.02
Noncluster Sample	-0.81 ± 0.02	-0.65 ± 0.04	-0.86 ± 0.03

Table 2
Radio Source Properties

Field	Redshift (z)	Opening Angle ($^{\circ}$)	Total Projected Physical Size (kpc)	Total Projected Physical Area (Mpc 2)	Radio Lobe Lengths	
					Lobe 1 ($''$)	Lobe 2 ($''$)
Cluster Candidate Sample						
COBRA072805.2+312857 ^a	$1.75 \pm 0.1^{b\ c}$	156.0 ± 2.3	471^{+26}_{-26}	$0.064^{+0.005}_{-0.005}$	$28.5^{+3.0}_{-3.0}$	$27.2^{+3.0}_{-3.0}$
COBRA074410.9+274011 ^a	$1.30 \pm 0.1^{b\ d}$	119.7 ± 3.6	256^{+25}_{-25}	$0.029^{+0.004}_{-0.004}$	$14.6^{+3.0}_{-3.0}$	$15.9^{+3.0}_{-3.0}$
COBRA100745.5+580713 ^a	$0.656 \pm 0.0003^{b\ c\ d\ e\ f}$	125.5 ± 3.1	271^{+21}_{-21}	$0.063^{+0.004}_{-0.003}$	$19.6^{+3.0}_{-3.0}$	$19.5^{+3.0}_{-3.0}$
COBRA100841.7+372513	$1.20 \pm 0.1^{c\ d}$	120.3 ± 2.1	412^{+34}_{-48}	$0.077^{+0.006}_{-0.006}$	$26.9^{+3.5}_{-5.7}$	$22.7^{+3.1}_{-3.1}$
COBRA103434.2+310352 ^a	$1.20 \pm 0.1^{b\ c}$	83.6 ± 0.2	675^{+26}_{-27}	$0.141^{+0.005}_{-0.006}$	$41.1^{+3.0}_{-3.0}$	$40.4^{+3.0}_{-3.0}$
COBRA104254.8+290719 ^a	1.35 ± 0.1^b	82.8 ± 0.2	520^{+26}_{-26}	$0.119^{+0.006}_{-0.006}$	$31.5^{+3.0}_{-3.0}$	$31.2^{+3.0}_{-3.0}$
COBRA113733.8+300010 ^a	$0.96 \pm 0.005^{b\ c\ e\ g}$	63.9 ± 2.6	335^{+25}_{-25}	$0.038^{+0.004}_{-0.004}$	$19.3^{+3.0}_{-3.0}$	$23.0^{+3.1}_{-3.1}$
COBRA121128.5+505253 ^b	1.364 ± 0.00042^f	150.4 ± 2.4	526^{+25}_{-25}	$0.217^{+0.006}_{-0.006}$	$32.5^{+3.0}_{-3.0}$	$30.1^{+3.0}_{-3.0}$
COBRA123940.7+280828 ^a	$0.92 \pm 0.0682^{b\ e}$	132.8 ± 1.0	581^{+26}_{-27}	$0.089^{+0.007}_{-0.007}$	$41.1^{+3.0}_{-3.0}$	$33.1^{+3.0}_{-3.0}$
COBRA130729.2+274659 ^b	1.144 ± 0.006^f	157.3 ± 3.6	724^{+25}_{-25}	$0.146^{+0.006}_{-0.005}$	$52.1^{+3.0}_{-3.0}$	$35.9^{+3.0}_{-3.0}$
COBRA135136.2+543955	$0.55 \pm 0.1^{b\ d}$	131.4 ± 1.8	479^{+43}_{-52}	$0.089^{+0.015}_{-0.018}$	$39.6^{+3.0}_{-3.0}$	$35.2^{+3.0}_{-3.0}$
COBRA135838.1+384722	$0.81 \pm 0.1^{b\ c}$	134.7 ± 5.1	257^{+25}_{-26}	$0.040^{+0.005}_{-0.005}$	$16.0^{+3.0}_{-3.0}$	$18.1^{+3.0}_{-3.0}$
COBRA141155.2+341510 ^b	1.818 ± 0.00019^f	141.3 ± 3.6	431^{+26}_{-26}	$0.087^{+0.004}_{-0.004}$	$28.7^{+3.0}_{-3.0}$	$22.5^{+3.0}_{-3.0}$
COBRA145023.3+340123	$1.20 \pm 0.1^{b\ d}$	107.2 ± 1.0	390^{+26}_{-26}	$0.059^{+0.005}_{-0.005}$	$23.0^{+3.0}_{-3.0}$	$24.1^{+3.0}_{-3.0}$
COBRA152647.5+554859 ^a	$1.10 \pm 0.1^{b\ c}$	159.9 ± 1.1	694^{+27}_{-28}	$0.229^{+0.011}_{-0.012}$	$37.4^{+3.0}_{-3.0}$	$47.6^{+3.0}_{-3.0}$
COBRA154638.3+364420 ^b	0.939 ± 0.00036^f	155.4 ± 1.0	643^{+25}_{-25}	$0.141^{+0.007}_{-0.007}$	$44.6^{+3.1}_{-3.1}$	$37.0^{+3.0}_{-3.0}$
COBRA164611.2+512915	$0.351 \pm 0.00008^{b\ c\ d\ f}$	48.2 ± 1.0	554^{+17}_{-16}	$0.050^{+0.004}_{-0.003}$	$53.5^{+3.0}_{-3.0}$	$58.5^{+3.2}_{-3.1}$
COBRA164951.6+310818	$0.52 \pm 0.07875^{b\ e}$	35.0 ± 0.4	412^{+35}_{-40}	$0.057^{+0.009}_{-0.010}$	$39.4^{+3.0}_{-3.0}$	$26.8^{+3.0}_{-3.0}$
COBRA170105.4+360958 ^a	0.80 ± 0.1^b	98.5 ± 0.6	290^{+25}_{-27}	$0.042^{+0.005}_{-0.005}$	$22.1^{+3.0}_{-3.0}$	$26.8^{+3.0}_{-3.0}$
COBRA171330.9+423502 ^a	$0.698 \pm 0.00018^{b\ f}$	136.4 ± 1.7	423^{+22}_{-22}	$0.107^{+0.007}_{-0.010}$	$30.7^{+3.1}_{-3.1}$	$32.0^{+3.0}_{-3.1}$
Noncluster Sample						
COBRA090102.7+420746 ^b	1.621 ± 0.00112^f	144.8 ± 0.1	294^{+26}_{-26}	$0.052^{+0.004}_{-0.004}$	$16.0^{+3.0}_{-3.0}$	$18.2^{+3.0}_{-3.0}$
COBRA090745.5+382740 ^b	1.743 ± 0.00017^f	145.6 ± 0.5	406^{+33}_{-26}	$0.112^{+0.005}_{-0.005}$	$21.2^{+3.0}_{-3.0}$	$26.7^{+3.7}_{-3.0}$
COBRA093726.6+365550	1.40 ± 0.1^b	140.6 ± 2.5	770^{+27}_{-27}	$0.139^{+0.008}_{-0.008}$	$51.3^{+3.0}_{-3.0}$	$40.1^{+3.1}_{-3.1}$
COBRA104641.5+282028	1.75 ± 0.1^b	54.8 ± 0.7	473^{+27}_{-27}	$0.027^{+0.006}_{-0.005}$	$33.1^{+3.1}_{-3.1}$	$22.9^{+3.0}_{-3.0}$
COBRA111707.3+305307	0.33 ± 0.0514^e	61.1 ± 0.33	317^{+34}_{-38}	$0.029^{+0.007}_{-0.006}$	$32.9^{+3.0}_{-3.0}$	$33.9^{+3.0}_{-3.0}$
COBRA120654.6+290742	$0.853 \pm 0.1139^{b\ c\ d\ e}$	178.7 ± 11.2	309^{+26}_{-28}	$0.053^{+0.005}_{-0.006}$	$22.4^{+3.0}_{-5.6}$	$17.9^{+3.0}_{-5.2}$
COBRA123347.0+354133	0.87 ± 0.1^b	132.2 ± 0.5	304^{+26}_{-26}	$0.047^{+0.005}_{-0.005}$	$20.6^{+3.0}_{-3.0}$	$18.7^{+3.0}_{-3.0}$
COBRA131854.0+231153	1.45 ± 0.1^b	153.2 ± 1.8	475^{+26}_{-26}	$0.027^{+0.005}_{-0.005}$	$35.4^{+3.0}_{-3.0}$	$20.8^{+3.0}_{-3.0}$
COBRA132903.2+253110 ^b	0.987 ± 0.00019^f	97.3 ± 1.2	543^{+27}_{-28}	$0.106^{+0.006}_{-0.005}$	$45.9^{+3.2}_{-3.4}$	$22.1^{+3.0}_{-3.0}$
COBRA143817.6+491233 ^b	1.358 ± 0.00094^f	113.4 ± 0.7	702^{+26}_{-26}	$0.157^{+0.012}_{-0.009}$	$41.4^{+3.0}_{-3.0}$	$42.1^{+3.0}_{-3.0}$
COBRA144207.1+562522	1.80 ± 0.1^b	137.9 ± 3.4	275^{+26}_{-26}	$0.067^{+0.005}_{-0.005}$	$16.0^{+3.0}_{-3.0}$	$16.5^{+3.0}_{-3.0}$
COBRA145656.0+501748	$0.88 \pm 0.1^{b\ c}$	140.9 ± 1.1	357^{+26}_{-27}	$0.062^{+0.005}_{-0.006}$	$23.1^{+3.0}_{-3.0}$	$23.0^{+3.0}_{-3.0}$
COBRA153317.4+391804 ^b	0.789 ± 0.00016^f	155.2 ± 1.2	838^{+23}_{-23}	$0.143^{+0.009}_{-0.008}$	$50.8^{+3.0}_{-3.0}$	$61.3^{+3.0}_{-3.0}$
COBRA155000.5+294953 ^b	2.328 ± 0.00062^f	109.7 ± 4.0	308^{+25}_{-25}	$0.049^{+0.004}_{-0.004}$	$17.7^{+3.0}_{-3.0}$	$19.9^{+3.0}_{-3.0}$
COBRA170443.9+295246	1.25 ± 0.1^f	140.5 ± 2.6	426^{+26}_{-26}	$0.071^{+0.005}_{-0.005}$	$29.4^{+3.0}_{-3.0}$	$21.7^{+3.0}_{-3.0}$

Notes.^a Bent radio AGNs hosted by BCGs identified in GM21.^b Photometric redshift estimate from comparing the $i - [3.6]$ color to EzGal models.^c Photometric redshift estimates from comparing the $[3.6] - [4.5]$ color to EzGal models.^d Photometric redshift estimates from comparing the $r - i$ color to EzGal models.^e Photometric redshift estimates from SDSS.^f Spectroscopic redshift from SDSS.^g Spectroscopic redshift from Blanton et al. (2003).^h Bent radio AGNs that are SDSS-identified quasars.

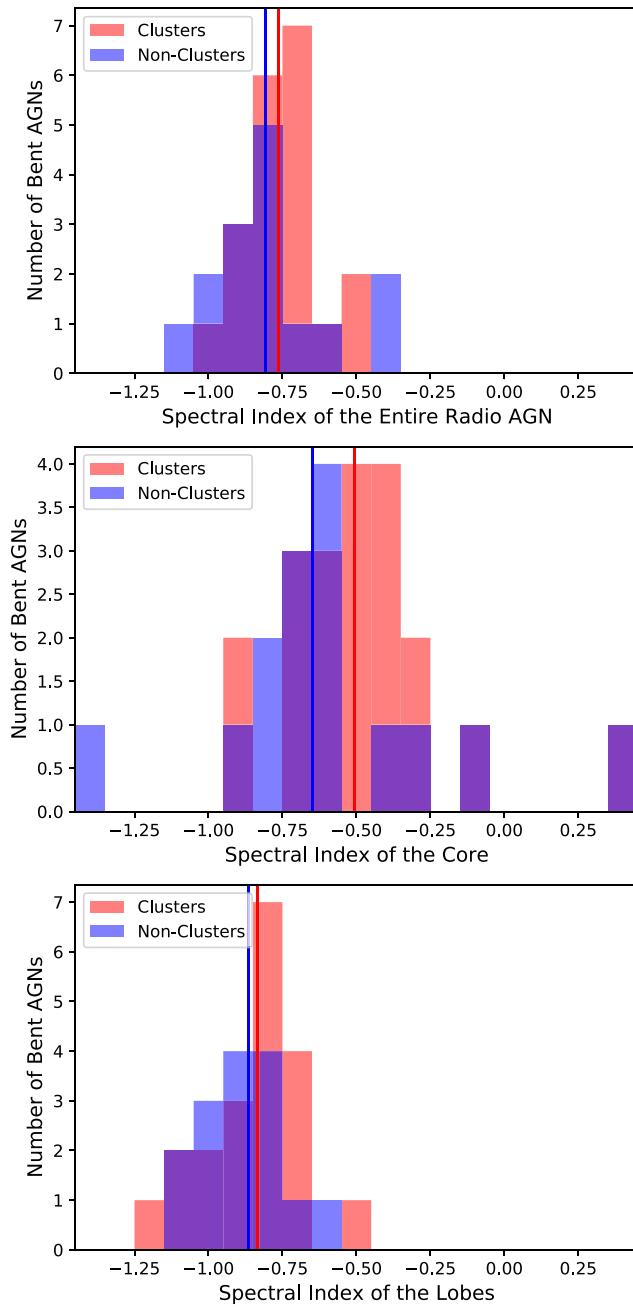


Figure 12. Histograms of the spectral index values of the entire radio AGN, the radio core, and the radio lobes. The radio AGNs in the cluster sample are shown in red and the radio AGNs in the noncluster sample are shown in blue. The median value of each distribution is shown as the vertical red and blue lines. While the total number of radio AGNs in each sample is 20 and 15, respectively, there are sources in each sample where the core region is smaller than the beam size. While we have included them here, they are flagged in Table 3.

of 20 bent radio AGNs in clusters with a flatter total spectral index (≥ -0.75) as compared to 4 of 15 in the noncluster sample.

Because quasars are more energetic than the typical bent radio AGNs in our sample, we also examine these populations separately. Using the four cluster sample quasars and the six noncluster sample quasars (see Table 3), we find a median spectral index of -0.82 ± 0.05 (clusters) and -0.80 ± 0.07 (nonclusters), which are in agreement with one another. In contrast, for the nonquasar sample, we find a median spectral

index of -0.76 ± 0.01 for the cluster sample and -0.83 ± 0.02 for the noncluster sample. This $> 2\sigma$ difference between the nonquasars would appear to further highlight a potential environmental impact on bent AGNs.

However, as the typical values of the spectral index of extended radio AGNs is between -0.7 and -0.8 (e.g., Kellermann & Owen 1988; Sarazin 1988; Condon 1992; Peterson 1997; Lin & Mohr 2007; Miley & De Breuck 2008), and the mean value of the spectral index of all cross-matched sources from LoTSS DR2 discussed in Shimwell et al. (2022) is -0.792 , all of our bent radio AGNs appear to be ordinary extended radio sources. While the overwhelming majority of sources can be described as having a spectral index with a flatter core and steeper lobes, even among sources where the core has a similar value to the lobes, we do have two sources, one in each sample (COBRA135838.1+384722 and COBRA170443.9+295246), where the core is steeper than the lobes (see Figures 20 and 21).

3.4.2. The Spectral Index of the Core and Lobes

To further probe the differences in the spectral index of bent AGNs as a function of environment, we measure the spectral index of the radio core. For this analysis, we fiducially estimate the area of each core based on the spectral index maps and plot a histogram of the mean values in the middle panel of Figure 12 (see Table 3). Although we report all values of the core spectral index, we flag measurements where either the core region is much smaller than the beam size, or in one case (COBRA104641.5+282028—which has the most positive spectral index core of any source in the noncluster sample) where the core is below the 10σ LoTSS threshold, but detected at the 3σ level in FIRST. For the sources where the core is a distinct separate component and smaller than the beam, we may be underestimating the error in these values.

We again find a difference between the populations of bent radio AGNs (see Table 1). The median core spectral index in the cluster sample is -0.51 ± 0.03 and in the noncluster sample is -0.65 ± 0.04 . While both values are flatter than the median of the total spectral index, we see an $\approx 2\sigma$ separation in these values. This is further enhanced in examining the distribution of core spectral indices. Specifically, 13 of 20 (65%) of bent radio sources in the cluster sample have mean core spectral indices ≥ -0.6 (11 of 17 [64.7%] if we remove those sources with a core that is poorly constrained), compared to 4 of 15 (26.7%) of bent radio AGNs in the noncluster sample (3 of 12 [25%] bent radio AGNs in the noncluster sample if we remove those with a poorly identified core). Although the difference is not great enough to claim a distinct separate sample without additional observations, we note that while not all bent AGNs in clusters have flatter cores, it does appear that sources with a flatter spectral core are more commonly in clusters.

Interestingly, of the SDSS-identified quasars in each sample (see Table 3), three of four quasars in the cluster sample fall into the flat-core subsample, with a median value of -0.39 ± 0.15 . However, in the noncluster sample, only three of six quasars fall within this sample, with a median value of -0.53 ± 0.06 . When probing the remaining nonquasar AGNs with well-defined cores, we find an even greater difference, with the median of the core spectral index of the cluster sample being -0.54 ± 0.03 and the median core spectral index of the noncluster sample being -0.75 ± 0.03 , which corresponds to a $> 3\sigma$ difference. Given the larger difference among the

Table 3
Radio Source Properties

Field	RMS		Flux Density			Radio Luminosity (10^{25} W Hz $^{-1}$)	Spectral Index		
	LoTSS (mJy beam $^{-1}$)	FIRST (mJy beam $^{-1}$)	LoTSS 10 σ (mJy)	LoTSS 3 σ (mJy)	FIRST (mJy)		Total	Core	Lobes
Cluster Candidate Sample									
COBRA072805.2+312857	0.20	0.15	115.8 \pm 11.6	80.2 \pm 8.0	17.0 \pm 1.7	30.8 $^{+7.2}_{-7.1}$	-0.67 \pm 0.14	-0.10 \pm 0.14	-0.96 \pm 0.15
COBRA074410.9+274011	0.10	0.14	36.9 \pm 3.7	18.8 \pm 1.9	2.0 \pm 0.3	3.6 $^{+0.9}_{-1.0}$	-0.98 \pm 0.16	-0.71 \pm 0.18 ^a	-1.07 \pm 0.17
COBRA100745.5+580713	0.06	0.15	67.6 \pm 6.8	52.6 \pm 5.3	11.2 \pm 1.2	4.1 $^{+0.5}_{-0.5}$	-0.67 \pm 0.14	-0.46 \pm 0.19 ^a	-0.68 \pm 0.15
COBRA100841.7+372513	0.07	0.17	96.5 \pm 9.6	65.2 \pm 0.65	14.6 \pm 1.5	14.6 $^{+3.9}_{-3.9}$	-0.65 \pm 0.14	-0.45 \pm 0.14	-0.84 \pm 0.15
COBRA103434.2+310352	0.08	0.14	2089.9 \pm 209.0	2010.1 \pm 201.0	315.1 \pm 31.5	397.7 $^{+107.1}_{-105.8}$	-0.80 \pm 0.14	-0.31 \pm 0.15	-0.81 \pm 0.14
COBRA104254.8+290719	0.10	0.14	137.6 \pm 13.8	83.9 \pm 8.4	13.3 \pm 1.4	16.7 $^{+4.4}_{-4.4}$	-0.80 \pm 0.14	-0.62 \pm 0.15	-0.84 \pm 0.14
COBRA113733.8+300010	0.08	0.14	52.3 \pm 5.2	46.3 \pm 4.6	8.3 \pm 0.9	6.7 $^{+0.9}_{-0.9}$	-0.74 \pm 0.14	-0.60 \pm 0.16	-0.76 \pm 0.14
COBRA121128.5+505253 ^b	0.09	0.12	1972.3 \pm 197.3	1850.2 \pm 185.0	234.5 \pm 23.4	407.0 $^{+64.2}_{-64.2}$	-0.90 \pm 0.14	-0.89 \pm 0.14	-0.89 \pm 0.14
COBRA123940.7+280828	0.10	0.15	120.6 \pm 12.1	61.4 \pm 6.1	10.1 \pm 1.1	8.2 $^{+2.0}_{-1.9}$	-0.78 \pm 0.14	-0.54 \pm 0.14	-0.79 \pm 0.15
COBRA130729.2+274659 ^b	0.08	0.16	333.6 \pm 33.3	187.7 \pm 18.8	22.6 \pm 2.3	31.7 $^{+4.7}_{-4.7}$	-0.92 \pm 0.14	-0.42 \pm 0.14	-1.2 \pm 0.15
COBRA135136.2+543955	0.05	0.16	111.9 \pm 11.2	70.6 \pm 7.1	13.0 \pm 1.3	4.0 $^{+1.9}_{-1.8}$	-0.73 \pm 0.14	-0.63 \pm 0.19 ^a	-0.74 \pm 0.14
COBRA135838.1+384722	0.08	0.11	89.7 \pm 9.0	80.6 \pm 8.1	12.3 \pm 1.2	8.5 $^{+3.0}_{-2.9}$	-0.82 \pm 0.14	-0.86 \pm 0.15	-0.81 \pm 0.14
COBRA141155.2+341510 ^b	0.13	0.14	1074.0 \pm 107.4	1036.8 \pm 103.7	190.6 \pm 19.1	393.0 $^{+69.7}_{-69.7}$	-0.73 \pm 0.14	-0.36 \pm 0.14	-0.97 \pm 0.14
COBRA145023.3+340123	0.07	0.14	77.8 \pm 7.8	52.7 \pm 5.3	7.3 \pm 0.8	10.0 $^{+2.7}_{-2.7}$	-0.86 \pm 0.15	-0.68 \pm 0.15	-0.94 \pm 0.15
COBRA152647.5+554859	0.09	0.16	563.6 \pm 56.4	484.2 \pm 48.4	78.8 \pm 7.9	85.0 $^{+24.0}_{-23.7}$	-0.79 \pm 0.14	-0.28 \pm 0.14	-0.83 \pm 0.14
COBRA154638.3+364420 ^b	0.11	0.14	394.1 \pm 39.4	308.0 \pm 30.8	89.7 \pm 9.0	49.8 $^{+6.8}_{-6.8}$	-0.54 \pm 0.14	0.35 \pm 0.14	-0.83 \pm 0.14
COBRA164611.2+512915	0.07	0.15	96.1 \pm 9.6	66.6 \pm 6.7	13.2 \pm 1.4	1.7 $^{+0.2}_{-0.2}$	-0.70 \pm 0.14	-0.48 \pm 0.15	-0.74 \pm 0.15
COBRA164951.6+310818	0.11	0.13	79.1 \pm 7.9	728.6 \pm 72.9	113.3 \pm 11.3	36.1 $^{+14.4}_{-13.9}$	-0.81 \pm 0.14	-0.54 \pm 0.14	-0.86 \pm 0.14
COBRA170105.4+360958	0.09	0.14	98.9 \pm 9.9	87.5 \pm 8.8	16.2 \pm 1.6	9.5 $^{+3.5}_{-3.5}$	-0.73 \pm 0.14	-0.70 \pm 0.14	-0.74 \pm 0.14
COBRA171330.9+423502	0.09	0.13	101.9 \pm 10.2	57.0 \pm 5.7	17.1 \pm 1.7	4.3 $^{+0.5}_{-0.5}$	-0.52 \pm 0.14	-0.44 \pm 0.17 ^a	-0.53 \pm 0.14
Noncluster Sample									
COBRA090102.7+420746 ^b	0.09	0.14	86.2 \pm 8.6	76.9 \pm 7.7	18.3 \pm 1.9	27.9 $^{+4.9}_{-4.7}$	-0.62 \pm 0.14	-0.44 \pm 0.14	-0.86 \pm 0.15
COBRA090745.5+382740 ^b	0.18	0.14	1047.9 \pm 104.8	971.4 \pm 97.1	160.3 \pm 16.0	331.1 $^{+51.1}_{-57.6}$	-0.78 \pm 0.14	-0.61 \pm 0.14	-0.87 \pm 0.14
COBRA093726.6+365550	0.10	0.12	149.8 \pm 15.0	54.8 \pm 5.5	7.3 \pm 0.8	12.8 $^{+2.9}_{-3.2}$	-0.87 \pm 0.14	-0.83 \pm 0.17	-0.88 \pm 0.15
COBRA104641.5+282028	0.11	0.13	18.4 \pm 1.8	8.3 \pm 0.8	3.3 \pm 0.4	4.2 $^{+0.9}_{-0.9}$	-0.41 \pm 0.15	0.36 \pm 0.22 ^c	-0.61 \pm 0.16
COBRA111707.3+305307	0.07	0.16	70.0 \pm 7.0	35.6 \pm 3.6	6.4 \pm 0.7	0.8 $^{+0.3}_{-0.3}$	-0.75 \pm 0.15	-0.65 \pm 0.16	-0.77 \pm 0.15
COBRA120654.6+290742	0.11	0.14	152.9 \pm 15.3	127.1 \pm 12.7	16.9 \pm 1.7	14.1 $^{+5.4}_{-5.2}$	-0.88 \pm 0.14	-0.68 \pm 0.14	-0.97 \pm 0.14
COBRA123347.0+354133	0.09	0.13	120.6 \pm 12.1	97.6 \pm 9.8	10.6 \pm 1.1	10.6 $^{+3.6}_{-3.5}$	-0.96 \pm 0.14	-0.79 \pm 0.14	-1.08 \pm 0.15
COBRA131854.0+231153	0.20	0.13	36.1 \pm 3.6	29.2 \pm 3.0	4.3 \pm 0.5	7.4 $^{+1.8}_{-1.8}$	-0.83 \pm 0.15	-0.71 \pm 0.20 ^a	-0.85 \pm 0.15
COBRA132903.2+253110 ^b	0.12	0.16	513.6 \pm 51.4	437.2 \pm 43.7	67.7 \pm 6.8	63.9 $^{+8.7}_{-8.9}$	-0.81 \pm 0.14	-0.61 \pm 0.14	-0.86 \pm 0.14
COBRA143817.6+491233 ^b	0.06	0.16	169.1 \pm 16.9	114.0 \pm 11.4	13.7 \pm 1.4	24.5 $^{+3.6}_{-3.9}$	-0.92 \pm 0.14	-0.34 \pm 0.15	-0.98 \pm 0.14
COBRA144207.1+562522	0.07	0.12	80.4 \pm 8.0	72.3 \pm 7.2	11.6 \pm 1.2	25.4 $^{+4.9}_{-5.9}$	-0.80 \pm 0.14	-0.71 \pm 0.14	-0.83 \pm 0.14
COBRA145656.0+501748	0.06	0.11	62.2 \pm 6.2	51.1 \pm 5.1	8.7 \pm 0.9	6.4 $^{+2.1}_{-2.0}$	-0.77 \pm 0.15	-0.65 \pm 0.16 ^a	-0.78 \pm 0.14
COBRA153317.4+391804 ^b	0.08	0.15	114.9 \pm 11.4	41.0 \pm 4.1	16.5 \pm 1.7	5.3 $^{+0.9}_{-0.7}$	-0.39 \pm 0.14	-0.11 \pm 0.14	-0.68 \pm 0.15
COBRA155000.5+294953 ^b	0.12	0.12	228.2 \pm 22.8	206.3 \pm 20.6	19.8 \pm 2.0	76.4 $^{+11.5}_{-15.1}$	-1.02 \pm 0.14	-0.85 \pm 0.14	-1.05 \pm 0.14
COBRA170443.9+295246	0.14	0.16	222.7 \pm 22.2	124.9 \pm 12.5	9.4 \pm 1.0	20.3 $^{+3.4}_{-5.5}$	-1.12 \pm 0.14	-1.35 \pm 0.16	-1.05 \pm 0.15

Notes.^a Cores where the spectral index is less well-constrained because it is much smaller than the beam size^b Bent Radio AGNs that are SDSS-identified quasars.^c Cores where the spectral index is less well-constrained because the core was not detected at the 10 σ level in LoTSS DR2

nonquasars and the small difference within the quasars, a flatter core spectral index in a nonquasar radio AGN could hint at further differences that might be used to identify radio AGNs in clusters (see Section 4.3 for a more complete discussion on the differences in the spectral index).

We further probe the differences in the populations of radio AGNs by examining the median spectral index of the radio lobes (see the bottom panel of Figure 12). The median values of the spectral index of the lobes slightly differ (-0.83 ± 0.02 for the cluster sample and -0.86 ± 0.03 for the noncluster sample, an $\approx 0.6\sigma$ offset). In general, both values are in agreement that the spectral index steepens as a function of the distance from the radio core.

4. Discussion

As highlighted in Sections 3.1, 3.2, and 3.4, we see slight differences in the properties of bent radio AGNs inside and outside of red sequence cluster candidates. To contextualize what these differences might mean in terms of the environment, we examine how some characteristics scale with cluster richness and compare our results to similar studies of the impact of the cluster environment on radio AGNs. We explore the impact of the cluster environment on radio source size in Section 4.1, the differences in radio luminosity in Section 4.2, and our spectral index of bent radio AGNs findings in Section 4.3.

4.1. Characterizing the Radio Lobes as a Function of Environment

4.1.1. Radio Source Asymmetry

There has been much work done to probe the impact of the ICM on radio lobe asymmetry. Rodman et al. (2019) used a sample of low- z ($z < 0.3$) FRI and FRII radio sources from the Radio Galaxy Zoo to probe the asymmetry of radio AGNs as a function of local environment and found an anticorrelation between the size of the radio lobe and the density of the surrounding environment. Yates-Jones et al. (2021) similarly probed this result using simulations and found that the density of the environment can be a driving factor in the asymmetry of radio AGNs. Garon et al. (2019), using a large sample of low- z ($z < 0.7$) bent radio AGNs found similar results using both the position of the bent radio AGN within the cluster and the pressure gradient across the cluster.

While a dearth of ICM studies of high- z COBRA clusters (only five have pointed X-ray observations from either Chandra or XMM-Newton; see E. L. Blanton et al. 2023, in preparation) prevents a detailed analysis of the density of the surrounding environment for most of our sample, Figure 9 shows that the most asymmetric sources are not in clusters. However, we do find nine bent sources in the cluster sample that are slightly more asymmetric (with a ratio of the area of the lobes < 0.8). We posit that this slight difference in radio lobe symmetry may result from the density of the surrounding gas.

For our sample, we analyze the potential role of the ICM density on radio lobe asymmetry using the combined overdensity measured in GM19 as our proxy for cluster richness (and ICM density; see Figure 13). We find a weak correlation with no evidence to reject the null hypothesis ($r_s = 0.266$, $p = 0.256$). While we see a stronger correlation between the degree of symmetry and overdensity if we focus on the more statistically robust $m^*+1 i - [3.6]$ cluster candidate sample

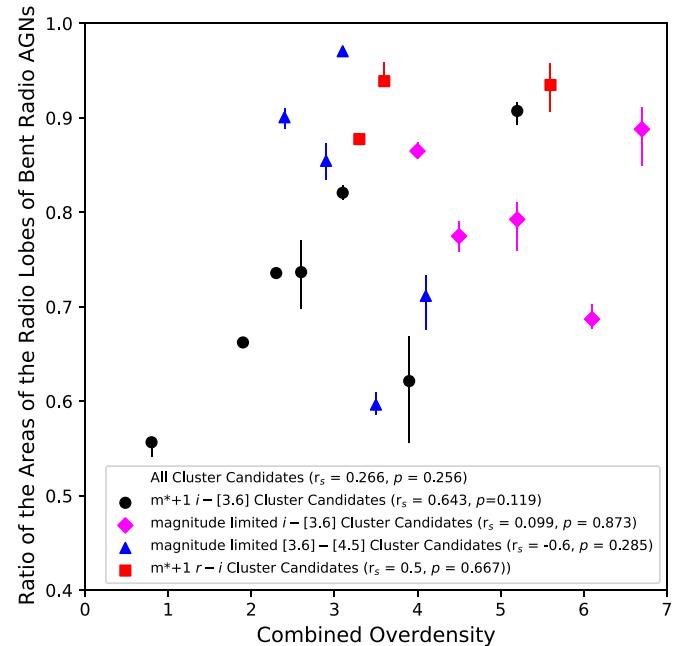


Figure 13. The ratio of the projected physical areas of the radio lobes as a function of the significance of the combined overdensity from GM19, our cluster richness proxy, for the cluster candidates. We plot the $m^*+1 i - [3.6]$ cluster candidates (the most statistically robust sample in the initial analysis) in black circles, the magnitude-limited $i - [3.6]$ cluster candidates in pink diamonds, the magnitude-limited $[3.6] - [4.5]$ cluster candidates in blue triangles, and the $m^*+1 r - i$ cluster candidates in red squares. We find a weak correlation with little evidence to reject the null hypothesis for the entire sample ($r_s = 0.266$, $p = 0.256$). While we find a stronger correlation in the $m^*+1 i - [3.6]$ sample, we again find no evidence to reject the null hypothesis, making this trend statistically improbable ($r_s = 0.643$, $p = 0.119$).

($r_s = 0.643$ and $p = 0.119$; see Section 2.1 and GM19 and GM21 for a full discussion of the COBRA subsamples), we again find no evidence to reject the null hypothesis. As some of our systems are in lower-mass clusters/groups, these systems may be more impacted by cluster weather and instabilities, which would make cluster richness a poor proxy for ICM density and may explain the more scattered distribution.

We examine the symmetry of bent radio AGNs in the noncluster sample as a function of environment and find no trend between the degree of lobe asymmetry and overdensity. Because the overdensity of many of these sources is poorly constrained or below the minimum three red sequence galaxy detection threshold in GM19, we do not show the correlation here.

To further study the environmental drivers of radio lobe asymmetry, we plot the degree of asymmetry as a function of the offset from the cluster center for the cluster candidate sample in Figure 14. Assuming a typical ICM density profile, for sources near the cluster center, any direction that the lobes expand into will encounter the same density gradient. For bent sources at greater offsets from the cluster center, the direction the radio lobes expand may result in greater differences due to the surrounding ICM density, particularly if both lobes do not point radially toward or away from the cluster center. Based on this simple model, differences in the surrounding ICM density distribution may cause some of the asymmetric lobes in our sample at large offsets (> 600 kpc), including for some of the BCGs found at larger offsets in our sample (see GM21 for a full discussion of the location of bent radio AGNs in COBRA

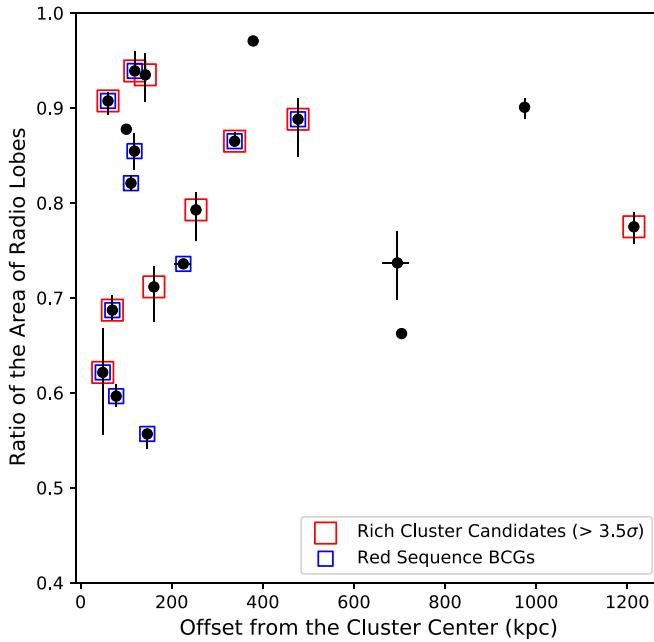


Figure 14. The ratio of the projected physical areas of the radio lobes as a function of the offset from the cluster center. The offset from the cluster centers is measured in GM19 and presented in GM21 and is based on the surface density of red sequence galaxies. Richer cluster candidates, with a combined overdensity greater than 3.5σ , are shown in red boxes. Bent radio AGNs hosted by BCGs are shown in blue boxes. We find no overall trend between the symmetry of radio lobes and the offset from the cluster center.

clusters and Gozaliasl et al. 2019 and Zhang et al. 2019 for further discussion of BCGs offset from the cluster center).

A number of our bent radio AGNs are located near the cluster center, including some of the most asymmetric sources. This result may be a function of the small sample size being probed, limiting our ability to identify highly symmetric sources at larger offsets from clusters. However, because we do not know the velocity differences between the galaxies (i.e., how fast the host galaxy is moving relative to the parent cluster), we propose that the bent radio AGNs near the cluster center are likely not all hosted by infalling galaxies, but rather hosted by brightest cluster galaxies (BCGs). As wide-angle-tail (WAT) radio sources are commonly hosted by BCGs (Owen & Rudnick 1976), the ICM is set in motion from a large-scale cluster-cluster merger (both major and minor mergers), which results in the bending of these radio tails (e.g., Blanton et al. 2011; Douglass et al. 2011; Paterno-Mahler et al. 2013; O’Dea & Baum 2023). As shown in the radio contours presented in Paterno-Mahler et al. (2013), bent radio AGNs associated with such systems can have asymmetries in their tails (highlighted by the width and extent of the tail in the bent radio AGN in A2029). Since 11 of 20 bent radio radio AGNs in the cluster sample are hosted by BCGs (see Figure 14) and an additional four are quasars, which are not identified as BCGs in GM21 because they are not on the red sequence, it is possible that gas sloshing yields the range of radio lobe asymmetries observed among bent radio AGNs near the cluster center.

Apart from gas sloshing, buoyant forces can also impact the bending of WATs (e.g., Gull & Northover 1973; Sakelliou et al. 1996; Smolčić et al. 2007). Since buoyancy is most impactful at larger radii (Smolčić et al. 2007), bending the sources toward regions of lower density, this may also explain

the asymmetry within the lobes. Furthermore, given the differences in the locations of the bent radio AGNs in clusters, the velocity of the host galaxy relative to the ICM may also impact radio lobe asymmetry, especially among the lower-mass, potentially less dynamically relaxed systems.

4.1.2. Radio Source Size

Beyond the radio source asymmetry, we also look at the difference in the size and location of these sources. We find a similar range of projected physical areas and lengths of bent AGNs in both samples, in agreement with Wing & Blanton (2011) at low redshift. However, much work has been done to examine populations of rich clusters at $z \approx 1$ hosting radio AGNs from either the MaDCoWS sample—done using FIRST observations to identify the radio sources and high-resolution VLA imaging in either the L band centered on 1.4 GHz with an angular resolution of $1.^{\prime}3$ or the C band centered on 5.5 GHz with an angular resolution of $1.^{\prime}0$ (e.g., Moravec et al. 2019, 2020)—and the ORELSE sample—done using 1.4 GHz JVLA observations with a beam size of $5.^{\prime}$ down to a $\approx 10 \mu\text{Jy}$ 1σ sensitivity (e.g., Shen et al. 2017, 2019, 2020). These studies found that radio AGNs farther from the cluster center are larger. If we assume spherically symmetric ICM distributions, these results imply that a less-dense ICM at greater cluster-centric radii allows for the formation of larger radio AGNs.

Following the normalization of the projected physical size discussed in Moravec et al. (2019), we do a similar analysis for our bent radio AGNs in clusters. As our measurements of the projected physical size are less well constrained, we measure the normalized projected physical area as a function of the offset from the cluster center (see Figure 15). To determine whether our data follow the trend in Moravec et al. (2019, 2020), we plot the slope of that trend re-normalized to our data (the gray line in Figure 15). As in GM21, we see no agreement with this trend and a large degree of scatter among our own sample.

While the sample of 49 radio AGNs in Moravec et al. (2020) includes some bent radio AGNs, the majority of the sources are straight (four are explicitly classified as bent tails, and one source that is in this sample is identified as an FRII). As such, it is possible that the more complicated geometries and projection effects associated with bent radio AGNs induce noise in our measurements and limit the effectiveness of this trend between AGN offset and AGN size. Additionally, although the cluster samples used in Moravec et al. (2020) and Shen et al. (2020), both of which show agreement with the trend, each have one cluster that overlaps with this COBRA sample, the typical cluster mass of both a MaDCoWS cluster and an ORELSE cluster is much higher than the expected cluster mass of a high- z COBRA target. The median MaDCoWS M_{500} is $\approx 10^{14.20} M_{\odot}$ (Gonzalez et al. 2019) and the median ORELSE cluster virial mass in Shen et al. (2020) is $10^{14.5} M_{\odot}$. However, for COBRA, only two of our clusters have mass estimates. COBRA113733.8 +300010 ($z = 0.96$), the only spectroscopically confirmed cluster (Blanton et al. 2003), has a virial mass of $\approx 10^{14.1} M_{\odot}$ (Lemaux et al. 2019), and COBRA164611.2+512915, the lowest-redshift cluster in our sample ($z = 0.351$), has a red-MaPPer richness of 42.5, corresponding to $M_{200} \approx 10^{14.23} M_{\odot}$ (Rykoff et al. 2014; Simet et al. 2017).

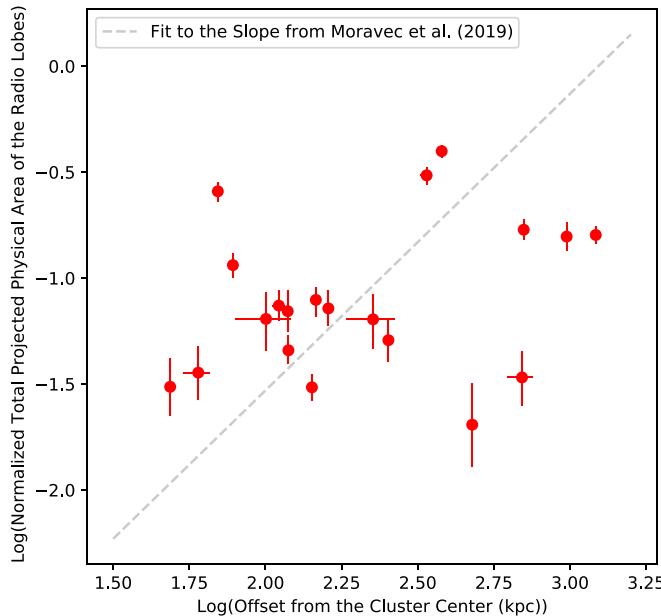


Figure 15. The log of the normalized projected physical area as a function of the log of the offset from the cluster center. The offsets from the cluster center are reported in GM21. The normalization of the fit from Moravec et al. (2019) is shown in the gray dashed line.

4.1.3. Opening Angle

The opening angle is thought to depend on the local environment. Thus, we plot the ratio of the areas of the radio lobes as a function of the size of the LoTSS opening angle of each bent radio AGN in both samples (Figure 16) to further explore the role of environment. Although we find a larger fraction of narrower sources in clusters, we find a strong similarity between the two samples and see no trend between opening angle and asymmetry. Moreover, Figure 16 shows a large population of wide AGNs (opening angle $> 135^\circ$) with moderate asymmetry (< 0.8) in both samples. Thus, while there is evidence that the opening angle is dependent on the density of the ICM (e.g., Garon et al. 2019), we find no evidence to conclude that the narrowness of the opening angle and the degree of lobe asymmetry are correlated.

It is interesting that we find that the widest sources in both environments in Figure 16 show a range of lobe asymmetries. In the cluster sample, this likely results from some bent sources being far from the cluster center, while others are near the cluster center, creating discrepancies in the local environment (see Figure 14). Similarly, we find a range of asymmetries among the narrowest sources in both samples (opening angles $< 80^\circ$), including the most asymmetric source. As we only have five narrow sources, we cannot determine whether this is driven by the environment, or the inclusion of both central and noncentral AGNs (for the three in the cluster sample). For the cluster sample, one is at the cluster center (COBRA113733.8+30010, offset < 50 kpc, asymmetry ≈ 0.62), while the other two (COBRA164611.2+512915 and COBRA164951.6+310818) are infalling near the cluster center. Given that the infalling sources are in clusters of different richness and differing asymmetries, we are unable to determine whether the narrow opening angle, the richness of the cluster, or a combination of the two results in the asymmetry. Ultimately, because our sample only contains a few classical narrow-angle-

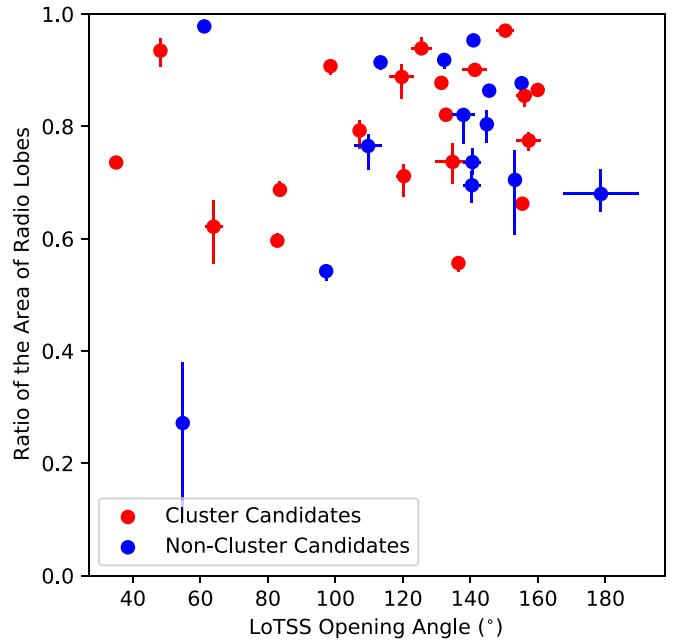


Figure 16. The ratio of the projected physical area of the radio lobes as a function of the LoTSS opening angle. While there is no clear distinction between the two populations, we find that narrower bent AGNs are generally more commonly in clusters, as expected.

tail (NAT) radio sources (opening angle $< 45^\circ$; e.g., Rudnick & Owen 1976; O'Dea & Owen 1985), a larger sample including more NATs, like the one in de Vos et al. (2021), is needed to determine the role a narrow opening angle plays in the asymmetry of the radio lobes.

4.1.4. Differences in Radio AGN Environment?

Despite the differences between the opening angles and spectral indices of bent radio AGNs inside and outside of clusters, we find much similarity between these two populations. It is possible that beyond the opening angle, the local environment has little impact on the remaining aspects of morphology. However, there are a few alternative possibilities specifically, with respect to our definition of the noncluster sample. While the majority of noncluster sources are in fields where GM19 was either unable to characterize a red sequence overdensity or had negative values, some systems are in weaker galaxy groups (overdensities of the order 1.5σ), which may imply that these systems are more like galaxy clusters than the field (i.e., the bent radio AGN is found either centrally within or on the outskirts of a quasi-symmetric intragroup medium). This hypothesis is strengthened if we look at the $3.6\ \mu\text{m}$ overdensities from Paterno-Mahler et al. (2017). For the fields in the noncluster sample, Paterno-Mahler et al. (2017) found 4 of 15 were in 2σ $3.6\ \mu\text{m}$ overdensities within a $1'$ search radius centered on the AGN and 9 of 15 were in 2σ $3.6\ \mu\text{m}$ overdensities within a $2'$ search radius. These larger-scale overdensities could point to structures offset from the AGNs. Additionally, 12 of 15 noncluster sources showed positive $3.6\ \mu\text{m}$ overdensities, which could further suggest that these sources reside in smaller structures, albeit ones not identified as red sequence cluster candidates in GM19. Moreover, despite the challenges in identifying high- z group candidates, there are numerous examples of bent radio AGNs found in galaxy groups (e.g., Ekers 1978; Venkatesan et al. 1994; Doe et al.

1995; Blanton et al. 2001; Freeland et al. 2008; Vardoulaki et al. 2021; Morris et al. 2022), which support this hypothesis.

Bent radio AGNs are not always found near the cluster center and in some cases can be found at large offsets. Sakelliou & Merrifield (2000) found one bent source offset by ≈ 1.4 Mpc from the cluster center in their sample of X-ray detected low- z clusters hosting bent AGNs. Edwards et al. (2010) identified the first known bent radio AGN in a filament offset from A1763 by 3.4 Mpc. Garon et al. (2019), in their study of the large-scale environment of bent radio AGNs at $z < 0.7$, found that more than 50% of their most bent sources are at offsets greater than $1.5 R_{500}$ (≈ 1 Mpc). While Garon et al. (2019) noted that bent sources at greater offsets are beyond the typical sphere of influence of a cluster, they found that for sources as offset as 2 Mpc, these bent sources still are in slightly overdense environments, implying that they reside in infalling groups or filaments. Similarly, de Vos et al. (2021) studied a sample of 208 NATs at $0.02 < z < 0.8$ and used photometric redshifts from Wen & Han (2015) to associate each NAT with its host cluster. Although most bent radio AGNs are closer to the cluster center, they also identified sources out to $10 R_{500}$ (which, following Garon et al. 2019, is ≈ 6.7 Mpc).

Given the $5' \times 5'$ FOV of Spitzer IRAC and the fact that all bent AGNs are at the center of the COBRA Spitzer imaging, GM19 could only probe between ≈ 0.74 Mpc ($2'/5$ at $z = 0.351$) and ≈ 1.27 Mpc ($2'/5$ at $z = 1.818$, the highest-redshift source in our cluster sample) from each AGN. For all of the clusters in GM19, the IRAC FOV limits our ability to identify bent AGNs at the largest offsets from the cluster center reported (e.g., Sakelliou & Merrifield 2000; Edwards et al. 2010; Garon et al. 2019; de Vos et al. 2021). Thus, it is possible our sample of noncluster bent radio AGNs is not a representative field sample, but rather a sample of bent radio AGNs infalling into clusters at distances beyond the scope of the analysis in GM19. This would imply that the differences seen in overdense environments (e.g., opening angle and asymmetry) are real, but could also account for the similarities, as these sources are associated with galaxy clusters just at varying distances.

4.2. The Impact of the Cluster Environment on Radio Luminosity

As shown in Figure 10, we find a strong correlation between the projected physical area and the radio luminosity of bent radio AGNs in COBRA cluster candidates. We see agreement between our result and that of Moravec et al. (2020), who found a similar trend between the projected physical extent and radio luminosity of radio AGNs in MaDCoWS clusters. Similarly, GM21 reported a strong correlation between the projected physical size and the radio luminosity of the larger high- z COBRA cluster candidate sample using FIRST observations.

We find differences between the cluster and noncluster sample in terms of the correlation between radio luminosity and the projected physical area, which leads to the question of whether environment impacts either property for bent radio AGNs. At low redshift, the literature has not reached a consensus. Croston et al. (2019) studied a large sample of low- z radio AGNs from LoTSS DR1 in clusters identified in redMaPPer, and found some evidence, with a large degree of scatter, that more powerful radio AGNs ($> 10^{24.5} \text{ WHz}^{-1}$) tend to reside in richer groups/clusters ($\approx 35 < \lambda < 45$; M_{200}

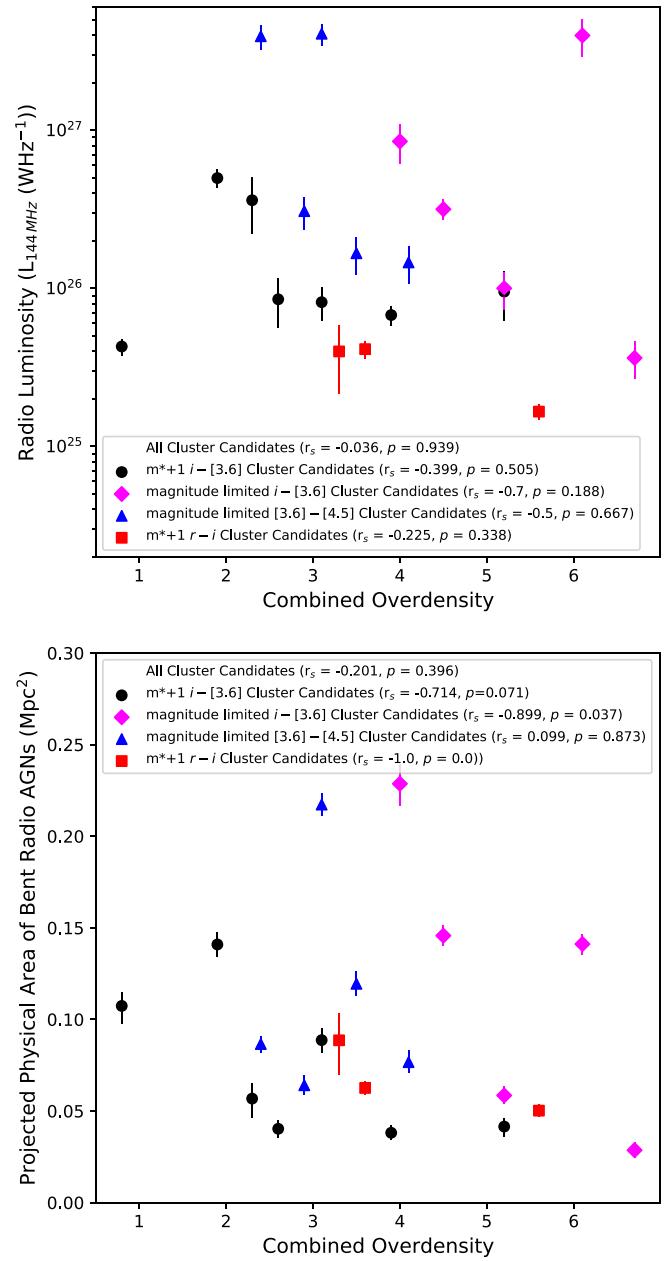


Figure 17. The radio luminosity ($L_{144 \text{ MHz}}$) as a function of the combined overdensity for the cluster candidate sample is shown in the top panel and the total projected physical area (Mpc^2) as a function of the combined overdensity is shown in the bottom panel. Each subsample is identified following the same legend as Figure 13. We find no evidence of any correlations between the combined overdensity and the radio luminosity. Although we see no trend between the projected physical size and combined overdensity for the entire sample, we do find a strong trend with weak to moderate evidence to reject the null hypothesis for the $m^{*}+1 i - [3.6]$, magnitude-limited $i - [3.6]$, and $m^{*}+1 r - i$ cluster samples.

$\approx 10^{14.14} - 10^{14.25} M_{\odot}$). In contrast, Wing & Blanton (2011) found no trends between richness and radio source size both inside and outside of clusters when looking at the parent sample of low- z bent AGNs used to create COBRA (although their overdensity measurements differ from redMaPPer's richness, many have been identified in redMaPPer and have masses between $10^{13.5}$ and $10^{14.5} M_{\odot}$).

We plot the radio luminosity as a function of the combined overdensity to determine the impact of the cluster environment for our LoTSS objects (see the top panel of Figure 17). While

we find some moderate trends between radio luminosity and richness within the individual subsamples, there is no statistical evidence to reject the null hypothesis. Furthermore, for the entire sample, we find no statistical evidence for trends between the richness of the cluster environment and the radio luminosity of our bent radio AGNs. Given that the trend identified by Croston et al. (2019) is strongest among systems likely more massive than many of our COBRA clusters and that they only probe a very narrow range of cluster masses, it is possible that the difference is due to the lower-mass cluster candidates hosting bent radio AGNs in our sample and in the Wing & Blanton (2011) sample or that the effect strongly depends on the radio luminosity of sources in similar cluster masses and is not an overarching trend.

To further determine if environment plays any role in the size/energetics of bent radio sources, we plot the total projected physical area as a function of the combined overdensity (lower panel of Figure 17). We find a weak trend with no evidence to reject the null hypothesis among the entire sample ($r_s = -0.201$, $p = 0.396$). Among the subsamples, we find moderate to strong anticorrelations with weak to moderate evidence to reject the null hypothesis among three of our cluster subsamples ($m^* + 1$ $i - [3.6] - r_s = -0.714$, $p = 0.071$; magnitude-limited $i - [3.6] - r_s = -0.899$, $p = 0.037$; $m^* + 1 - i - r_s = -1.0$, $p = 0.0$).

Given that our total sample agrees with the results from Wing & Blanton (2011), it is difficult to determine whether the trend of larger radio AGNs being found in poorer clusters is real, or due to small number statistics. However, if true, it would agree with some of the inferences from Moravec et al. (2019) and Moravec et al. (2020), where the size of the radio AGNs was a function of local environment (assuming weaker overdensities have a weaker ICM density). Additionally, given the difference in redshift between our sample and Wing & Blanton (2011), this could also represent redshift evolution. As shown in Vardoulaki et al. (2021), the ICM gets denser at lower redshift, which can result in narrower bent AGNs because of the additional space for jet interactions. Thus, a similar scenario could occur where at high redshift, the necessary ICM/gas density to constrict AGNs and bend the radio lobes only exists in richer clusters, while at low- z , clusters, groups, filaments, and/or fossil groups could have the minimum gas density to not only bend the radio lobes, but severely constrict the size. Additionally, due to the 120° selection criteria of the parent sample of COBRA from Wing & Blanton (2011), our sample is sensitive to physically larger radio AGNs at high- z , which may account for the correlations we identify.

Ultimately, as with our analysis of the asymmetry of bent radio AGN lobes, our analysis of the radio luminosity of bent AGNs finds little evidence of environmental dependence. While it is possible that the projected physical area and luminosity of radio AGNs are not strongly dependent on environment, it is also possible that the environments of the cluster and noncluster sample are similar. As we currently lack X-ray observations for the majority of these sources and require a larger sample to better determine if the trends we identify between cluster richness and projected physical area are real, we plan to address these questions in the future using new data from eRosita and continuing observations from LOFAR.

4.3. A Comparison of Spectral Indices

As discussed in Section 3.4, we measure the spectral index for the entire radio source, the spectral index of the radio core, and the spectral index of the radio lobes for the bent radio AGNs in our sample. We find that the median total spectral index of bent radio AGNs in clusters is slightly flatter than those not in clusters (-0.76 ± 0.01 versus -0.81 ± 0.02) and that this is especially true among the nonquasar populations (-0.76 ± 0.01 versus -0.83 ± 0.02). Past studies of bent AGNs or AGNs in clusters in general do not use identical frequencies, nor do they have the same sensitivity, so we can only make a qualitative comparison in terms of the similarities of the properties of bent radio AGNs and the environments of these sources.

In comparing our samples to other observations of bent AGNs (either head-tails or WATs), we find similar results to sources in low- z clusters (e.g., Patnaik et al. 1986; O'Donoghue et al. 1990; Di Gennaro et al. 2018a, 2018b; O'Dea & Baum 2023). O'Donoghue et al. (1990) presented one of the largest samples of low- z WATs with uniform observations that have spectral index measurements. Using VLA observations at 6 cm (≈ 4.996 GHz) and 20 cm (≈ 1.444 GHz), O'Donoghue et al. (1990) found that the cores of these bent sources tend to have spectral indices of $-0.6 < \alpha < -0.5$ and that the lobes are much steeper, with $\alpha \approx -1.5$, in agreement with our measurements. Using individual targets, both Gendron-Marsolais et al. (2020), who measured the spectral index between 344 MHz and 610 MHz, and Patnaik et al. (1986), who measured the spectral index between 1.444 GHz and 4.996 GHz, saw bent radio AGNs with a flatter radio core, typically of the order of $\alpha > -0.5$ and steeper radio lobes ($\alpha < -1$). Similarly, Di Gennaro et al. (2018b) analyzed the spectral index in a similar fashion to our method (but between 235 MHz and 8.4 GHz), and also found flatter radio cores and steeper diffuse regions, which are akin to our radio lobes, for their two tailed radio AGNs.

At high redshift ($1.5 < z < 3.2$), Barthel & Miley (1988) identified a sample of 30 bent radio quasars (opening angles $< 160^\circ$) among their sample of 80 high- z quasars with steeper spectral indices ($\alpha < -0.6$ measured using VLA observations at 1.4 GHz and 5 GHz). While we only study 10 quasars and our redshift range is $0.79 < z < 2.346$, we find similar results, with all but the two lowest redshift quasars having a total spectral index < -0.6 . Although Barthel & Miley (1988) did not identify which sources might be in protocluster/cluster environments, the similarities of our samples may be indicative of these quasars being found in regions of increased overdensity.

While not specifically studying bent radio AGNs, both Mahony et al. (2016) and de Gasperin et al. (2018) studied large samples of radio AGNs out to high- z . Mahony et al. (2016) used LOFAR observations (150 GMz) and 1.4 GHz observations to observe radio AGNs in the Lockman Hole and found that steep spectrum sources ($-1.2 < \alpha < -0.5$) make up 82.1% of the sources in their Lockman-Wide sample, while flat sources ($\alpha > -0.5$) make up 5.7% of sources. Similarly, we find all but two of our sources would be characterized as steep spectrum sources based on their criteria. de Gasperin et al. (2018) probed 147 MHz–1.4 GHz across 80% of the radio sky and found a weighted mean spectral index of -0.7870 ± 0.0003 . Given the similar spectral index measurements between our samples, Mahony et al. (2016), and de Gasperin et al. (2018), it appears that the spectral index of bent

radio AGNs is not a uniquely defining characteristic and that these sources have spectral indices of typical radio AGNs.

In regards to the total spectral index and the spectral index of the radio core, we find an environmental difference between those in clusters and those not in clusters. While the spectral index of the cores in both samples is generally flatter than the lobes, the spectral index of cores in the cluster sample are flatter than those in the noncluster sample with a 2σ significance. Similar to our study, Dabhade et al. (2020) measured the mean spectral index of giant radio galaxies in both clusters and the field. Although they used different surveys, they span a similar frequency range (NVSS at 1400 MHz and LoTSS DR1), which makes the results comparable. Dabhade et al. (2020) found a mean spectral index of -0.79 for the giant radio quasars and -0.78 for the giant radio galaxies. Although not the focus of their analysis, Dabhade et al. (2020) identified 20 giant radio galaxies in low- z ($z < 0.55$) clusters and reported that 14 have spectral indices flatter than or equal to their reported mean values (two of the sources have no measurement), which indicates a lower median spectral index for radio AGNs in clusters, in agreement with our findings.

To interpret the difference in the spectral index of the cores of bent radio AGNs, we look to de Gasperin et al. (2018), who fully modeled the spectral index as a function of the life cycle of the AGN. de Gasperin et al. (2018) found that younger, more core-dominated AGNs are both brighter and have flatter cores. As the AGN ages, the lobes become more dominant as the AGN jets shut down. Following this model for our sources would imply that the sources in clusters are more likely to be either younger or more recently turned on. Given the obvious environmental difference, we argue that the bent radio AGNs in clusters have flatter cores because material was more recently accreted onto them, which may be due to the denser environment. This could be particularly true of the quasars in the noncluster sample, which have flatter cores, though not as flat as those in the cluster sample, indicating that they might be slightly older than their cluster counterparts. Based on this interpretation, having a flatter spectral index, particularly in the core, would appear to be the strongest signpost for determining which future detected bent radio AGNs are in clusters without optical/IR follow-up (although not having a strong flat core does not rule out a bent radio AGN from being in a cluster).

5. Conclusion

This paper is the fourth paper in the series probing high- z bent radio AGNs identified as part of the high- z COBRA survey. This is the first to fully probe the impact of the cluster environment on the morphology of bent radio AGNs to determine differences between bent radio AGNs inside and outside of clusters. Using new publicly available observations from LoTSS DR2, we identify a sample of 20 high- z bent radio AGNs in clusters and 15 bent radio AGNs not in clusters. Using this sample, we examine differences between the projected physical size, projected physical area, radio source asymmetry, opening angle, radio luminosity, and for the first time for a large sample of high- z bent radio AGNs, the spectral index of these sources. Below is a summary of our findings.

1. *Differences in the projected physical size, area, and the asymmetry of bent radio AGNs.* For each of the 35 radio AGNs in our sample, we estimate the projected physical size and projected physical area of the bent radio AGNs within 10σ contours. Within the cluster sample, we find

tentative evidence of moderate to strong anticorrelations between the projected physical area of each radio AGN and the richness of the host cluster. However, we find no distinct differences between the sources in clusters and the sources not in clusters.

2. *Measurements of the opening angle.* We measure the opening angle of each bent AGN observed with LoTSS and find agreement with previous measurements from FIRST. For the cluster sample, we do not see a statistically robust trend between the size of the opening angle and cluster richness. However, we do find more narrow sources in clusters, implying some environmental differences in the populations of bent radio AGNs.
3. *Probing radio luminosity.* We estimate the radio luminosity and examine it as a function of projected physical area and environment. We find a moderate positive correlation between radio luminosity and the projected physical area of the AGN for the cluster sample and a weak correlation with no evidence to reject the null hypothesis for the noncluster sample. Despite these differences, a K-S test implies that both samples are drawn from the same sample, further highlighting their similarities.
4. *New measurements of the spectral index of bent radio AGNs.* We measure the spectral index for a large sample of bent radio AGNs at high- z . Using observations from LoTSS DR2 and FIRST, we find the median values of the spectral index differ by $\approx 1.6\sigma$ (-0.76 ± 0.01 and -0.81 ± 0.02 , respectively) and the core spectral index differs by $\approx 2\sigma$ for the total sample (-0.51 ± 0.03 and -0.65 ± 0.04) and $> 3\sigma$ for the nonquasars with a well-defined core (-0.54 ± 0.03 and -0.75 ± 0.03). We find that 13 of 20 cluster bent radio AGNs have a flatter core (≥ -0.6) as compared to 4 of 15 in the noncluster sample. As the flatness of the core indicates the age of the radio emission, this suggests that these sources are indeed in richer cluster environments with more recent accretion onto the supermassive black hole. It is possible that future cluster surveys using bent radio AGNs should target AGNs with a core with a flatter spectral index.

The goal of this study was to identify differences between the populations of bent radio AGNs based on their large-scale environment to improve the identification of high- z cluster candidates. Although we found that sources with narrower opening angles and flatter spectral cores are indeed more commonly found in clusters, we find a surprising degree of similarity between the samples. Ultimately, our populations of bent radio AGNs are similar in regards to their projected physical size, area, and radio luminosity, which may point to further similarities between the environments of these radio AGNs.

While the noncluster sample is not located in rich cluster overdensities of red sequence galaxies as shown in GM19, it is possible that these sources reside in bent radio AGNs that are outside the cluster center, possibly beyond the area probed in GM19, either in the outer cluster regions or falling along filaments. As bent radio AGNs require a gaseous medium to bend the radio lobes, the differences we find may be due to the differing affects of the local and large-scale environment (e.g., Rodman et al. 2019; Yates-Jones et al. 2021). This interpretation would explain why some of the radio source properties are similar (e.g., the projected physical area and the radio luminosity), despite the obvious outliers (narrow opening angles and flatter radio cores) that are indicative of richer cluster

environments. This stands to further strengthen the use of all bent radio AGNs as signposts for structure in large sky surveys.

Given the similarities between these two populations, we aim to continue to study populations of bent radio AGNs to better characterize the host environments of our noncluster sample. Although X-ray observations of distant clusters and filaments are observationally expensive, with more data becoming available from telescopes like eRosita, we will continue to better quantify the gaseous environments to improve constraints on the role of the local gas density on radio morphology, especially in regards to the asymmetry. In particular, we plan to better constrain the role of the opening angle on asymmetry by studying samples of NATs and straight radio AGNs. Given the abundance of data available at low- z , including measurements of the ambient density, we will further probe the environments of low- z bent radio AGNs to improve constraints on the mass of the host clusters and further characterize populations of bent radio AGNs that are potentially in filaments or fossil group galaxies.

Additionally, we plan to further constrain the energetics of radio AGNs by probing larger samples at lower redshifts, where we can more accurately map differences over the entire radio source. With new radio observations, from surveys like those being conducted with Meerkat and SKAO, we plan to better map the entirety of the spectral index of the radio source, allowing us to probe the asymmetry of the lobes of bent radio AGNs.

Lastly, we will continue to work to identify new bent radio AGNs out to the highest redshifts with future LoTSS data releases. Although none of the sources in our sample are characterized as ultra-steep sources ($\alpha < -1.4$; e.g., Mahony et al. 2016; de Gasperin et al. 2018), these radio AGNs tend to be associated with high- z ($z > 2$) protoclusters (e.g., Wylezalek et al. 2013, 2014) and include quasars with bent radio AGN morphology (Barthel & Miley 1988). As we continue to search for bent radio AGNs at progressively higher redshifts, it will be interesting to see if we identify any such sources in clusters, as they could be the signposts for some of the earliest forming dense ICMs.

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Facilities: LDT, Spitzer, Sloan, LOFAR, FIRST.

Software: astropy (Astropy Collaboration et al. 2013), CASA (McMullin et al. 2007).

Appendix A Identifying Host Galaxies

As mentioned in Section 2.1, Paterno-Mahler et al. (2017) used the 3.6 μm Spitzer IRAC images to identify host galaxies for each of the bent radio AGNs in the entire high- z COBRA survey. This was done by overlaying radio contours from FIRST on the Spitzer images and cross-matching with SDSS to avoid accidentally identifying an incorrect low- z host. In Figures 18 and 19, we have overlaid the LoTSS contours (the same contours as in Figures 2 and 3) on the 3.6 μm Spitzer IRAC images to show the location of the host galaxy relative to the radio source.

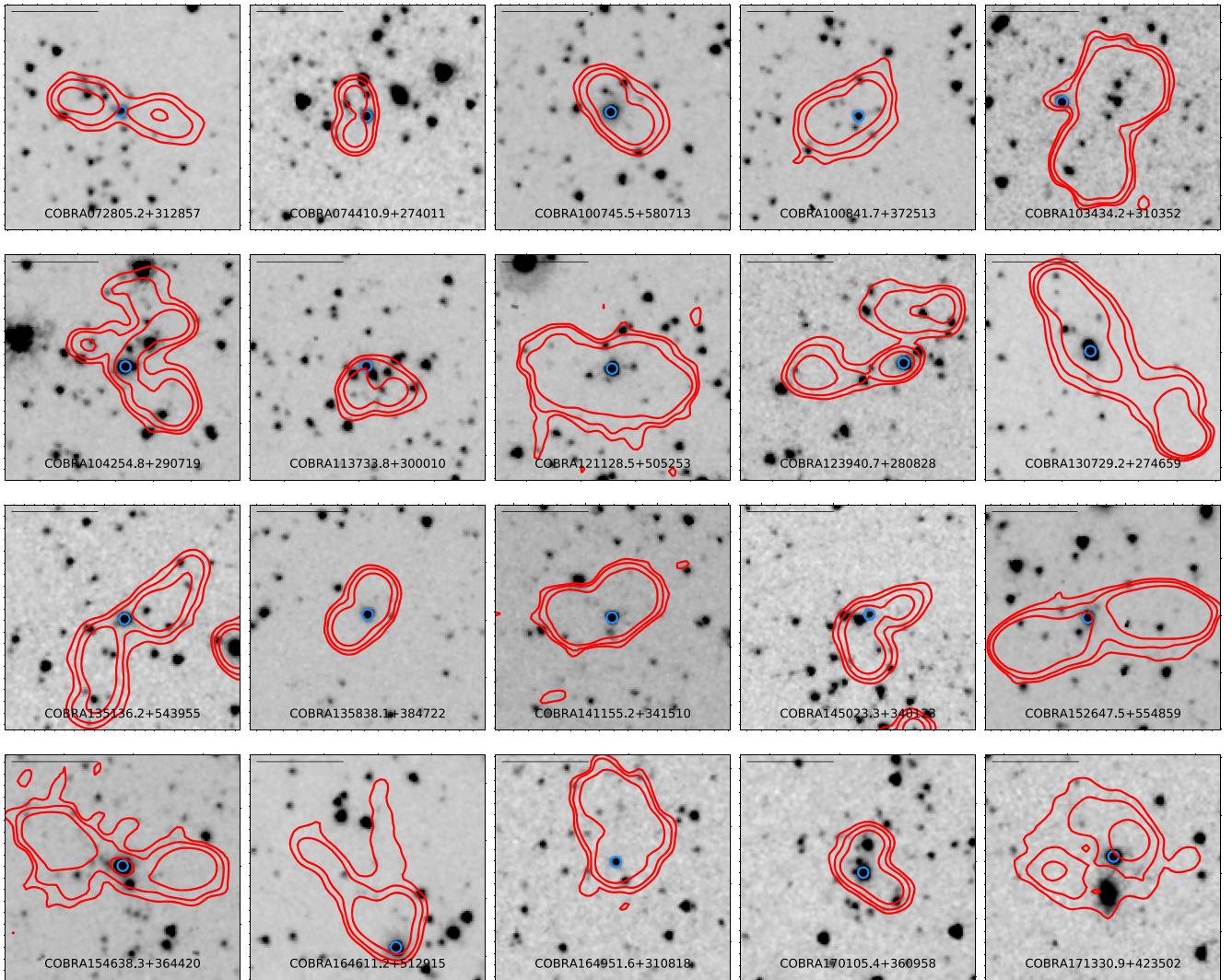


Figure 18. 3.6 μm Spitzer IRAC Cutouts of the 20 bent, double-lobed radio sources in the COBRA cluster candidate sample observed by LoTSS. Each image shows the same FOV as in Figure 2, and the three red contours show the same 10 σ , 20 σ , and 50 σ contours based on the measurement of the LoTSS rms noise in each image. Here, the blue circle identifies the host galaxy. The black line shows 0.5''.

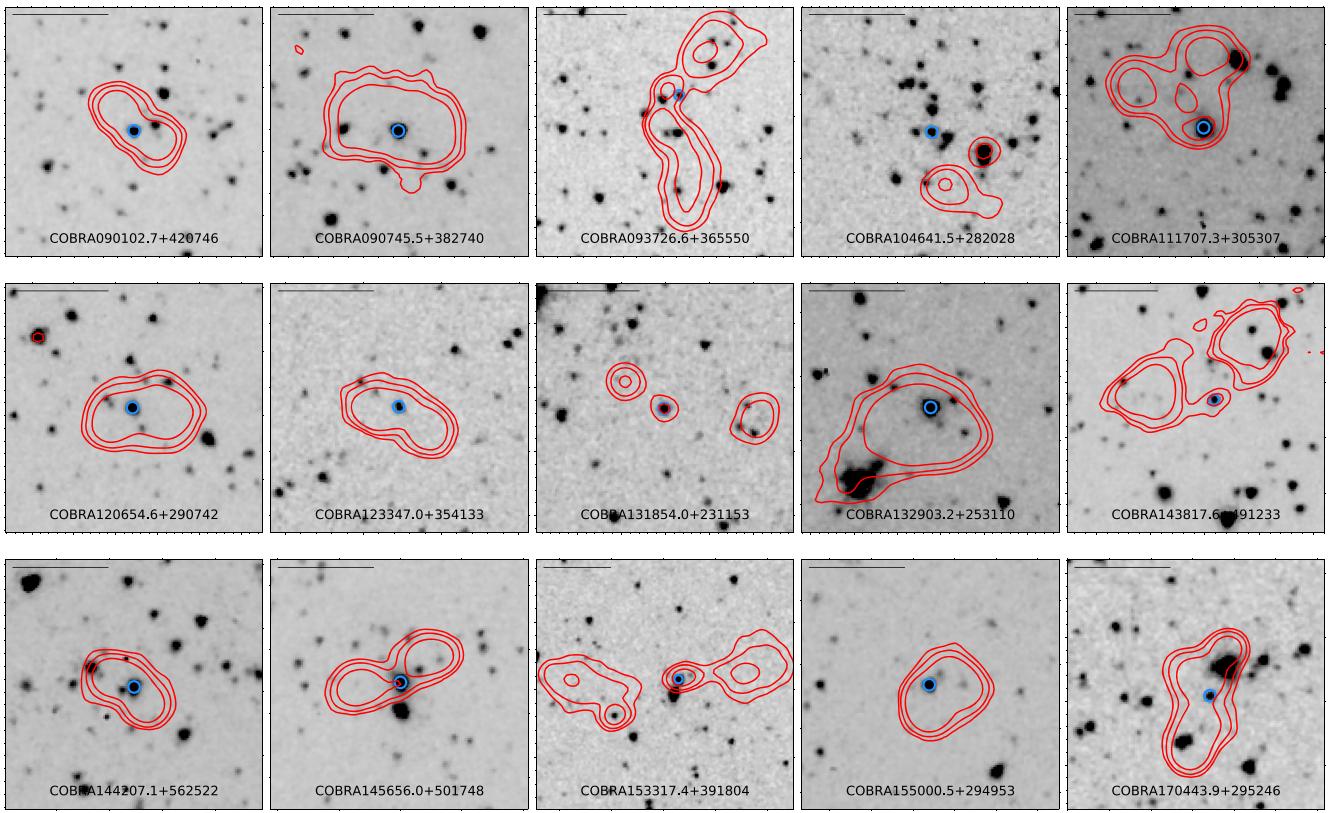


Figure 19. $3.6\text{ }\mu\text{m}$ Spitzer IRAC Cutouts of the 15 bent, double-lobed radio sources in the noncluster sample observed by LoTSS. Each image shows the same FOV as in Figure 3. The red contours show the same contours as in Figure 3, while the blue circle identifies the host galaxy.

Appendix B Spectral Index Maps

As mentioned in Section 2.3, because the LoTSS beam size is much larger than the pixel size, individual pixel measurements of the spectral index are highly correlated. Thus, we limit our analysis of the spectral index to regions of approximately the beam size or larger (e.g., the entire radio source, the radio core, or the radio lobes). However, to allow for an examination of the regions probed for each bent source, we do present our spectral index maps in Figures 20 and 21. As can be seen, we have a number of sources in the cluster sample, as well as a few in the noncluster sample, with very flat spectral index measurements in the core relative to the lobes. As seen in our reported values, we see a steepening of the spectral index away from the core.

Additionally, we plot the error in the pixel measurements of the spectral index to highlight the regions that are subject to a slightly higher degree of error in Figure 22 and 23. We measure the error following Equation (1) in Di Gennaro et al. (2018b). Unlike the spectral index maps shown in Figures 20 and 21, where the value does change rapidly across each AGN, we find a relatively uniform degree of error across the majority of each source, with only slightly higher errors at the edges of the sources. This is in agreement with Gendron-Marsolais et al. (2020), who showed that the inner regions of their radio AGNs have roughly constant error values, similar to our own. We do see a higher degree of error in the core of COBRA104641.5 +282028, likely because the core region was not detected within the 10σ LoTSS contours.



Figure 20. Spectral index maps of bent radio AGNs in the cluster sample. Each pixel corresponds to a $1''.5 \times 1''.5$ region of the sky. The $6''$ LoTSS beam is shown as a gray circle in the bottom-right corner of each image. The 3σ FIRST contours are shown. Although most images are shown at the same size ($\approx 0''.45 \times 0''.45$ centered on the radio core), a few of the images are scaled differently to encompass the entire radio AGN. We use a uniform color bar to show the value of the spectral index. We see examples of both core-dominated and lobe-dominated sources within our sample. Despite showing values of the spectral index for each individual pixel (the pixels are highly correlated because the pixel size is smaller than the beam size), we limit our analysis to larger structures and only identify the cores and lobes of our radio AGNs (see Table 3).

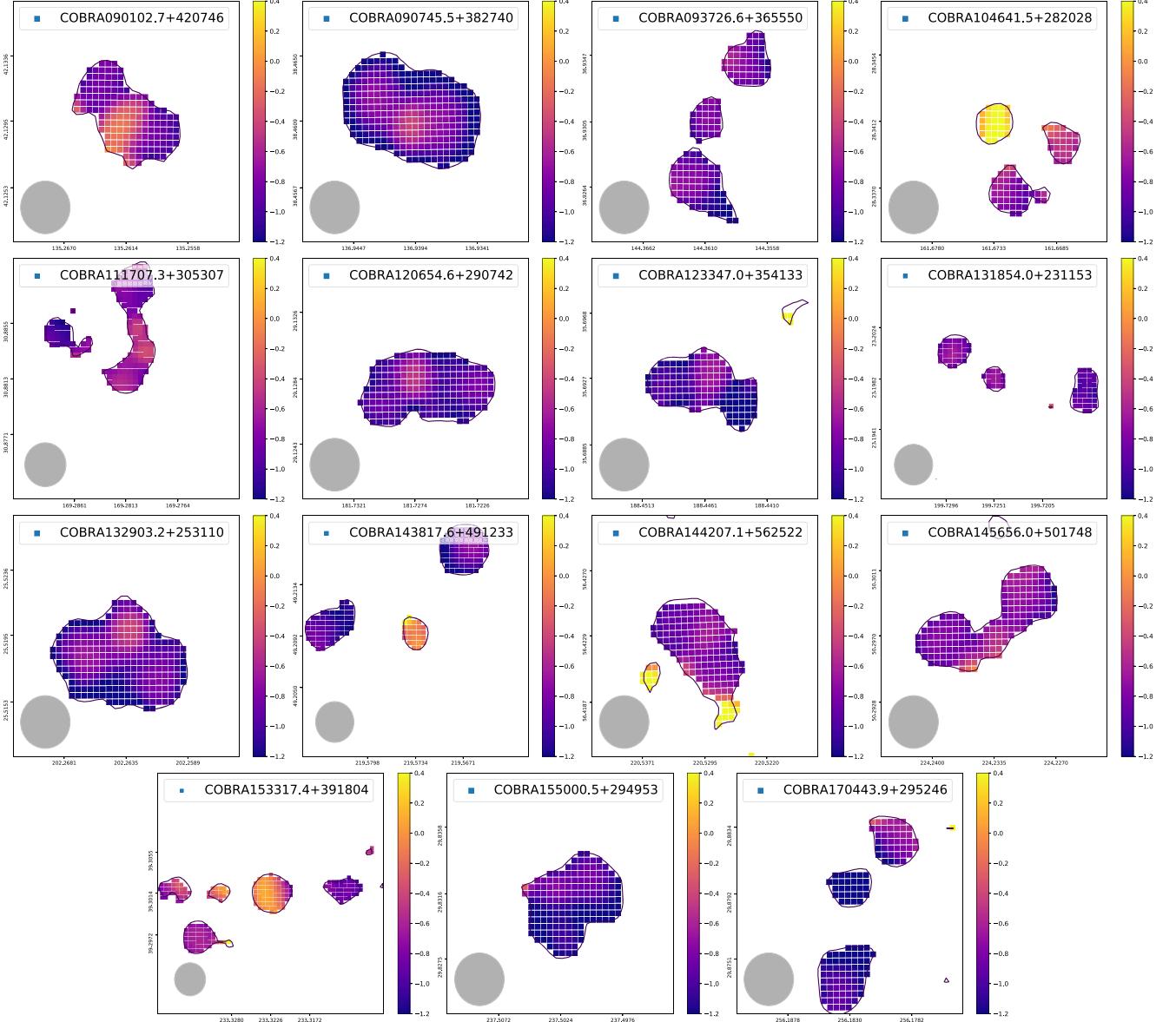


Figure 21. Spectral index maps of bent radio AGNs in the noncluster sample. As in Figure 20, each pixel corresponds to a $1.\overset{\circ}{5} \times 1.\overset{\circ}{5}$ region of the sky. Although most images are scaled the same ($\approx 0.\overset{\circ}{45} \times 0.\overset{\circ}{45}$ centered on the radio core), a few are scaled differently to encompass the entire radio AGN. We use the same color bar as in Figure 20.

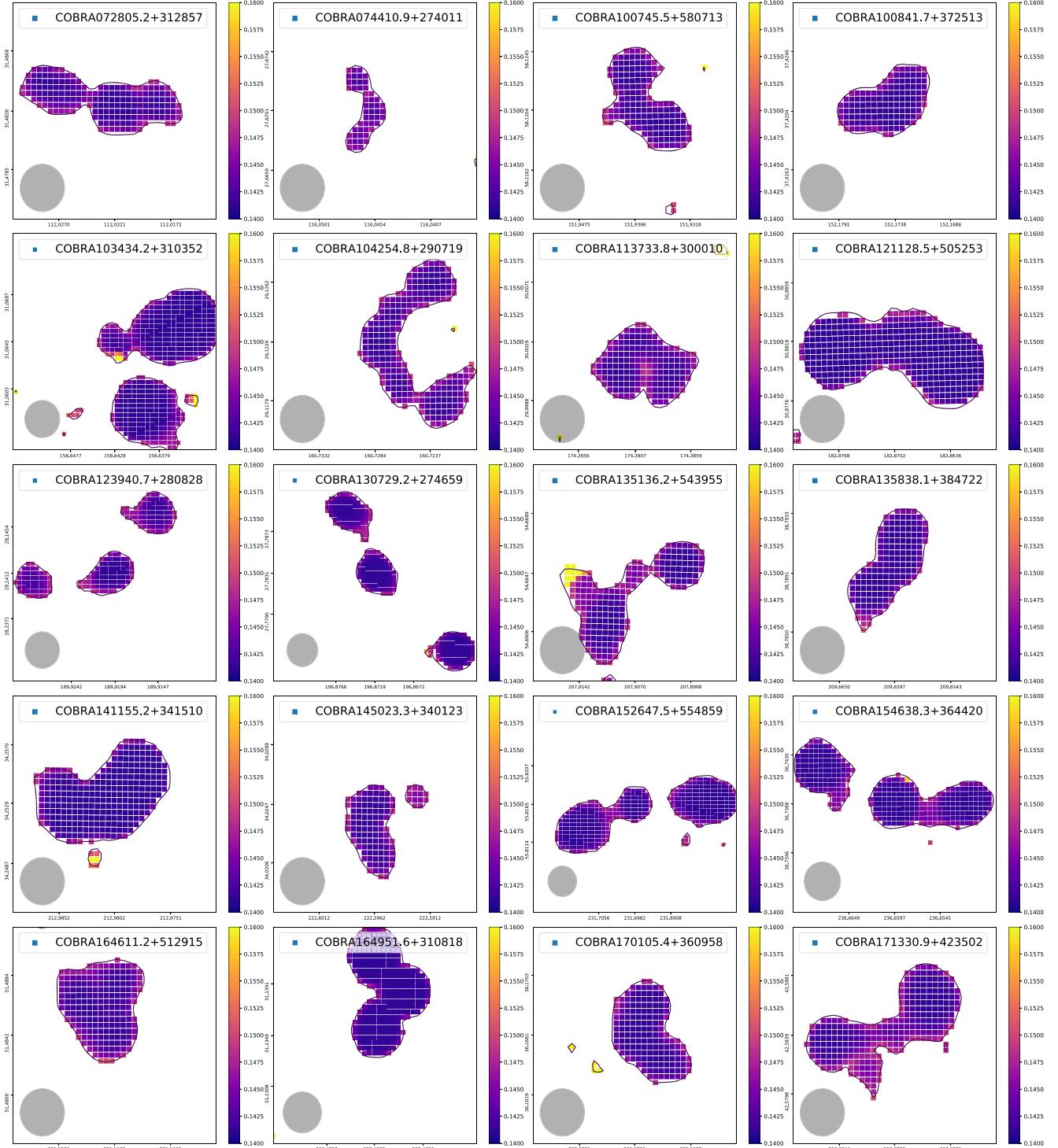


Figure 22. Spectral index maps of bent radio AGNs in the cluster sample. Each pixel corresponds to a $1.^{\circ}5 \times 1.^{\circ}5$ region of the sky. The $6''$ LoTSS beam is shown as a gray circle in the bottom-right corner of each image. The black contours show the 3σ FIRST contours. All images show the same region as in Figure 20. While we see a roughly uniform error, we do use a uniform color bar to show our measure of the error in the value of the spectral index. As with the spectral index maps, we again note that the pixels are highly correlated because the pixel size is smaller than the beam size.

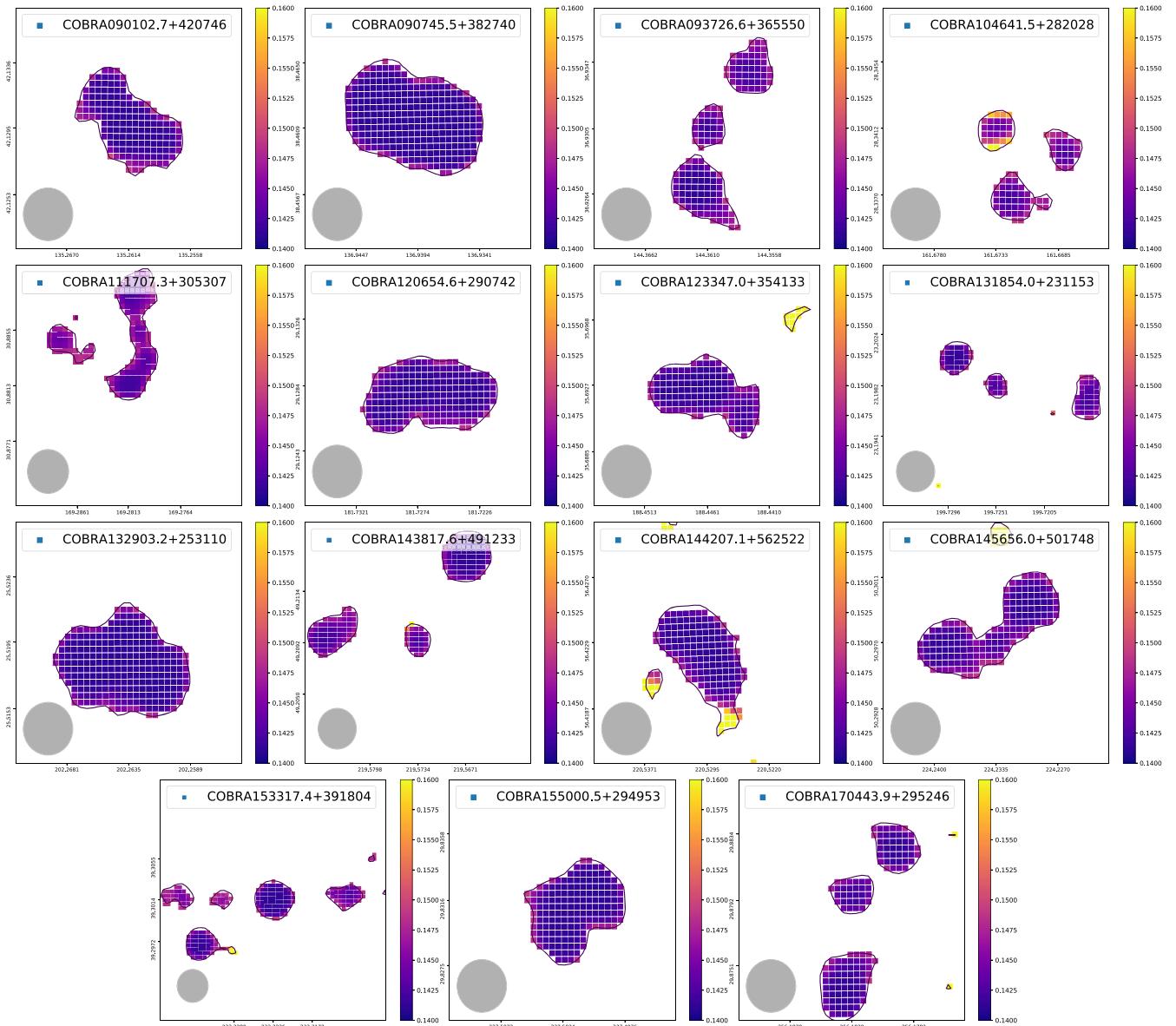


Figure 23. Spectral index error maps of bent radio AGNs in the noncluster sample. As in Figure 22, each pixel corresponds to a $1.\overset{\circ}{5} \times 1.\overset{\circ}{5}$ region of the sky, and the images cover the same region as those in Figure 21. We use the same color bar as in Figure 22.

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References

- Astropy Collaboration, Robitaille, T. P., Tollerud, E. J., et al. 2013, [A&A, 558, A33](#)
 Astropy Collaboration, Price-Whelan, A. M., Sipőcz, B. M., et al. 2018, [AJ, 156, 123](#)
 Barthel, P. D., & Miley, G. K. 1988, [Natur, 333, 319](#)
 Becker, R. H., White, R. L., & Helfand, D. J. 1995, [ApJ, 450, 559](#)
 Best, P. N. 2000, [MNRAS, 317, 720](#)
 Blakeslee, J. P., Franx, M., Postman, M., et al. 2003, [ApJ, 596, L143](#)
 Blanton, E. L., Gregg, M. D., Helfand, D. J., Becker, R. H., & Leighly, K. M. 2001, [AJ, 121, 2915](#)
 Blanton, E. L., Gregg, M. D., Helfand, D. J., Becker, R. H., & White, R. L. 2000, [ApJ, 531, 118](#)
 Blanton, E. L., Gregg, M. D., Helfand, D. J., Becker, R. H., & White, R. L. 2003, [AJ, 125, 1635](#)
 Blanton, E. L., Paterno-Mahler, R., Wing, J. D., et al. 2015, in IAU Symp. 313, Extragalactic Jets from Every Angle, ed. F. Massaro, [315](#)
 Blanton, E. L., Randall, S. W., Clarke, T. E., et al. 2011, [ApJ, 737, 99](#)
 Cai, Z., Fan, X., Bian, F., et al. 2017, [ApJ, 839, 131](#)

- Cerulo, P., Couch, W. J., Lidman, C., et al. 2016, *MNRAS*, **457**, 2209
- Condon, J. J. 1992, *ARA&A*, **30**, 575
- Cooke, E. A., Hatch, N. A., Rettura, A., et al. 2015, *MNRAS*, **452**, 2318
- Cooke, E. A., Hatch, N. A., Stern, D., et al. 2016, *ApJ*, **816**, 83
- Croston, J. H., Hardcastle, M. J., Mingo, B., et al. 2019, *A&A*, **622**, A10
- Dabholkar, P., Röttgering, H. J. A., Bagchi, J., et al. 2020, *A&A*, **635**, A5
- de Gasperin, F., Intema, H. T., & Frail, D. A. 2018, *MNRAS*, **474**, 5008
- de Vos, K., Hatch, N. A., Merrifield, M. R., & Mingo, B. 2021, *MNRAS*, **506**, L55
- Di Gennaro, G., van Weeren, R. J., Hoeft, M., et al. 2018a, *ApJ*, **865**, 24
- Di Gennaro, G., Venturi, T., Dallacasa, D., et al. 2018b, *A&A*, **620**, A25
- Doe, S. M., Ledlow, M. J., Burns, J. O., & White, R. A. 1995, *AJ*, **110**, 46
- Douglass, E. M., Blanton, E. L., Clarke, T. E., Randall, S. W., & Wing, J. D. 2011, *ApJ*, **743**, 199
- Edwards, L. O. V., Fadda, D., & Frayer, D. T. 2010, *ApJL*, **724**, L143
- Eisenhardt, P. R., De Propris, R., Gonzalez, A. H., et al. 2007, *ApJS*, **169**, 225
- Ekers, R. D. 1978, *A&A*, **69**, 253
- Fanaroff, B. L., & Riley, J. M. 1974, *MNRAS*, **167**, 31P
- Freeland, E., Cardoso, R. F., & Wilcots, E. 2008, *ApJ*, **685**, 858
- Fujita, Y., Kawakatu, N., & Shlosman, I. 2016, *PASJ*, **68**, 26
- Galametz, A., Stern, D., Eisenhardt, P. R. M., et al. 2009, *ApJ*, **694**, 1309
- Garon, A. F., Rudnick, L., Wong, O. I., et al. 2019, *AJ*, **157**, 126
- Gendre, M. A., Best, P. N., Wall, J. V., & Ker, L. M. 2013, *MNRAS*, **430**, 3086
- Gendron-Marsolais, M., Hlavacek-Larrondo, J., van Weeren, R. J., et al. 2020, *MNRAS*, **499**, 5791
- Golden-Marx, E., Blanton, E. L., Paterno-Mahler, R., et al. 2019, *ApJ*, **887**, 50
- Golden-Marx, E., Blanton, E. L., Paterno-Mahler, R., et al. 2021, *ApJ*, **907**, 65
- Gonzalez, A. H., Decker, B., Brodin, M., et al. 2015, *ApJL*, **812**, L40
- Gonzalez, A. H., Gettings, D. P., Brodin, M., et al. 2019, *ApJS*, **240**, 33
- Gozaliasl, G., Finoguenov, A., Tanaka, M., et al. 2019, *MNRAS*, **483**, 3545
- Gull, S. F., & Northover, K. J. E. 1973, *Natur*, **244**, 80
- Hardcastle, M. J., & Croston, J. H. 2020, *NewAR*, **88**, 101539
- Hardcastle, M. J., Sakelliou, I., & Worrall, D. M. 2005, *MNRAS*, **359**, 1007
- Hatch, N. A., Wylezalek, D., Kurk, J. D., et al. 2014, *MNRAS*, **445**, 280
- Hill, G. J., & Lilly, S. J. 1991, *ApJ*, **367**, 1
- Hung, D., Lemaux, B. C., Gal, R. R., et al. 2020, *MNRAS*, **491**, 5524
- Kellermann, K. I., & Owen, F. N. 1988, in Galactic and Extragalactic Radio Astronomy, Astronomy and Astrophysics Library, ed. K. I. Kellermann & G. L. Verschuur (New York: Springer), 563
- Ledlow, M. J., & Owen, F. N. 1996, *AJ*, **112**, 9
- Lemaux, B. C., Gal, R. R., Lubin, L. M., et al. 2012, *ApJ*, **745**, 106
- Lemaux, B. C., Tomczak, A. R., Lubin, L. M., et al. 2019, *MNRAS*, **490**, 1231
- Lin, Y.-T., & Mohr, J. J. 2007, *ApJS*, **170**, 71
- Longair, M. S., & Seldner, M. 1979, *MNRAS*, **189**, 433
- Lubin, L. M., Gal, R. R., Lemaux, B. C., Kocevski, D. D., & Squires, G. K. 2009, *AJ*, **137**, 4867
- Mahony, E. K., Morganti, R., Prandoni, I., et al. 2016, *MNRAS*, **463**, 2997
- Mancone, C. L., & Gonzalez, A. H. 2012, *PASP*, **124**, 606
- McMullin, J. P., Waters, B., Schiebel, D., Young, W., & Golap, K. 2007, in ASP Conf. Ser. 376, Astronomical Data Analysis Software and Systems XVI, ed. R. A. Shaw, F. Hill, & D. J. Bell (San Francisco, CA: ASP), 127
- Mei, S., Blakeslee, J. P., Stanford, S. A., et al. 2006, *ApJ*, **639**, 81
- Mei, S., Holden, B. P., Blakeslee, J. P., et al. 2009, *ApJ*, **690**, 42
- Miley, G., & De Breuck, C. 2008, *A&ARv*, **15**, 67
- Miller, C. J., Nichol, R. C., Reichart, D., et al. 2005, *AJ*, **130**, 968
- Miller, N. A., Owen, F. N., Burns, J. O., Ledlow, M. J., & Voges, W. 1999, *AJ*, **118**, 1988
- Mingo, B., Croston, J. H., Hardcastle, M. J., et al. 2019, *MNRAS*, **488**, 2701
- Minkowski, R. 1960, *ApJ*, **132**, 908
- Mo, W., Gonzalez, A., Brodin, M., et al. 2020, *ApJ*, **901**, 131
- Mo, W., Gonzalez, A., Stern, D., et al. 2018, *ApJ*, **869**, 131
- Moravec, E., Gonzalez, A. H., Stern, D., et al. 2019, *ApJ*, **871**, 186
- Moravec, E., Gonzalez, A. H., Stern, D., et al. 2020, *ApJ*, **888**, 74
- Morris, M. E., Wilcots, E., Hooper, E., & Heinz, S. 2022, *AJ*, **163**, 280
- Morsony, B. J., Miller, J. J., Heinz, S., et al. 2013, *MNRAS*, **431**, 781
- Muzzin, A., Wilson, G., Yee, H. K. C., et al. 2009, *ApJ*, **698**, 1934
- Noiro, G., Stern, D., Mei, S., et al. 2018, *ApJ*, **859**, 38
- Noiro, G., Vernet, J., De Breuck, C., et al. 2016, *ApJ*, **830**, 90
- O'Dea, C. P., & Baum, S. A. 2023, *Galax*, **11**, 67
- O'Dea, C. P., & Owen, F. N. 1985, *AJ*, **90**, 954
- O'Donoghue, A. A., Eilek, J. A., & Owen, F. N. 1993, *ApJ*, **408**, 428
- O'Donoghue, A. A., Owen, F. N., & Eilek, J. A. 1990, *ApJS*, **72**, 75
- Owen, F. N., & Rudnick, L. 1976, *ApJL*, **205**, L1
- Papovich, C. 2008, *ApJ*, **676**, 206
- Paterno-Mahler, R., Blanton, E. L., Brodin, M., et al. 2017, *ApJ*, **844**, 78
- Paterno-Mahler, R., Blanton, E. L., Randall, S. W., & Clarke, T. E. 2013, *ApJ*, **773**, 114
- Patnaik, A. R., Malkan, M. A., & Salter, C. J. 1986, *MNRAS*, **220**, 351
- Peterson, B. M. 1997, An Introduction to Active Galactic Nuclei (Cambridge: Cambridge Univ. Press)
- Prestage, R. M., & Peacock, J. A. 1988, *MNRAS*, **230**, 131
- Proctor, D. D. 2006, *ApJS*, **165**, 95
- Rodman, P. E., Turner, R. J., Shabala, S. S., et al. 2019, *MNRAS*, **482**, 5625
- Rudnick, L., & Owen, F. N. 1976, *ApJL*, **203**, L107
- Rykoff, E. S., Rozo, E., Busha, M. T., et al. 2014, *ApJ*, **785**, 104
- Sakelliou, I., & Merrifield, M. R. 2000, *MNRAS*, **311**, 649
- Sakelliou, I., Merrifield, M. R., & McHardy, I. M. 1996, *MNRAS*, **283**, 673
- Sarazin, C. L. 1988, X-ray Emission from Clusters of Galaxies, Cambridge Astrophysics Ser. (Cambridge: Cambridge Univ. Press)
- Shen, L., Lemaux, B. C., Lubin, L. M., et al. 2021, *ApJ*, **912**, 60
- Shen, L., Liu, G., Zhang, M.-F., et al. 2020, *ApJ*, **902**, 101
- Shen, L., Miller, N. A., Lemaux, B. C., et al. 2017, *MNRAS*, **472**, 998
- Shen, L., Tomczak, A. R., Lemaux, B. C., et al. 2019, *MNRAS*, **484**, 2433
- Shimwell, T. W., Hardcastle, M. J., Tasse, C., et al. 2022, *A&A*, **659**, A1
- Shimwell, T. W., Röttgering, H. J. A., Best, P. N., et al. 2017, *A&A*, **598**, A104
- Shimwell, T. W., Tasse, C., Hardcastle, M. J., et al. 2019, *A&A*, **622**, A1
- Silverstein, E. M., Anderson, M. E., & Bregman, J. N. 2018, *AJ*, **155**, 14
- Simet, M., McClintock, T., Mandelbaum, R., et al. 2017, *MNRAS*, **466**, 3103
- Smolčić, V., Schinnerer, E., Finoguenov, A., et al. 2007, *ApJS*, **172**, 295
- Snyder, G. F., Brodin, M., Mancone, C. M., et al. 2012, *ApJ*, **756**, 114
- Stocke, J. T., Perlman, E. S., Gioia, I. M., & Harvanek, M. 1999, *AJ*, **117**, 1967
- van Haarlem, M.P., Wise, M.W., & Gunst, A.W. 2013, *A&A*, **556**, A2
- Vardoulaki, E., Vazza, F., Jiménez-Andrade, E. F., et al. 2021, *Galax*, **9**, 93
- Venkatesan, T. C. A., Batuski, D. J., Hanisch, R. J., & Burns, J. O. 1994, *ApJ*, **436**, 67
- Wen, Z. L., & Han, J. L. 2015, *ApJ*, **807**, 178
- Wilson, G., Muzzin, A., Yee, H. K. C., et al. 2009, *ApJ*, **698**, 1943
- Wing, J. D., & Blanton, E. L. 2011, *AJ*, **141**, 88
- Wing, J. D., & Blanton, E. L. 2013, *ApJ*, **767**, 102
- Wylezalek, D., Galametz, A., Stern, D., et al. 2013, *ApJ*, **769**, 79
- Wylezalek, D., Vernet, J., De Breuck, C., et al. 2014, *ApJ*, **786**, 17
- Yates-Jones, P. M., Shabala, S. S., & Krause, M. G. H. 2021, *MNRAS*, **508**, 5239
- Zhang, Y., Jeltema, T., Hollowood, D. L., et al. 2019, *MNRAS*, **487**, 2578
- Zirbel, E. L. 1997, *ApJ*, **476**, 489